Correlation between Gamma-Ray bursts and Gravitational Waves

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The cosmological origin of γ-ray bursts (GRBs) is now commonly accepted and, according to several models for the central engine, GRB sources should also emit at the same time gravitational waves (GWBs). We have performed two correlation searches between the data of the resonant gravitational wave detector AURIGA and GRB arrival times collected in the BATSE 4B catalog. No correlation was found and an upper limit \( h_{\text{RMS}} \leq 1.5 \times 10^{-18} \) on the averaged amplitude of gravitational waves associated with γ-ray bursts has been set for the first time.

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I. INTRODUCTION

Thirty years after their discovery, Gamma Ray Bursts (GRBs) are still the most mysterious objects in the Universe, as their properties have not yet been associated to any well-known object, while for instance pulsars have been almost immediately identified as final objects of the evolution of stars and quasars have been well framed among galaxy nuclei in particular in the Seyfert class. Unfortunately “standard” astrophysical objects do not exist for GRBs. The “non standard” compact astrophysical objects which might reproduce the experimental characteristic of GRBs involve black-holes (BH), neutron stars (NS) and massive stars.

The wide range of characteristics prevents the systematic classification of GRBs. Their energy spectrum is continuous and non-thermal, and covers the range between 1 keV and 1 MeV; their emission lasts from 10 ms to \( 10^3 \) s. The Burst And Transient Source Experiment (BATSE) has detected GRBs with an event rate of about one per day between 1991 and 2000.

One of the most important results of observations is that many GRBs are at cosmological distances \( z \approx 1 \). This fact, recently confirmed also by the satellite BeppoSAX [2], which has detected some optical counterpart with redshifts close to \( z = 1 \) [3]. This implies the emission of a huge amount of electromagnetic radiation: in few seconds the GRBs release an energy of \( 10^{52} - 10^{54} \) erg. On the other hand, this implies that the GRBs are rare events in a galaxy: in fact, as the rate of GRB events is of the order \( 1/\text{day} \), their rate amounts to 1 event/\( 10^6 \) years per galaxy [7].

The theoretical framework which gives a coherent explanation of the experimental data set of GRBs is the so-called “fireball” model: an “inner engine” produces a flux of relativistic energy of the order \( 10^{52} \) erg and it is extremely compact and clearly not visible; this relativistic flux of particles is a kind of “fireball” able to produce electromagnetic radiation in a optically thin shell of matter. The inner engine works for 1-3 days, producing the afterglow, but the bulk of the explosion lasts typically 10 s (with peaks around 0.5 and 30 s and a variation of six orders of magnitude, \( 10^{-3} \div 10^3 \) s) when the major part of the energy is emitted.

The inner engine progenitors are likely to be coalesc- ing binary systems, such as NS-BH or NS-NS systems, or single stars that collapse into BH in supernova-like events (the so-called “collapsar” models [8]). If this is the correct explanation of the GRBs, then we would have compact astrophysical systems for which gravity plays an important role, and gravitational waves would be emitted [9–11].

This scenario has motivated us to investigate the behavior of the gravitational wave detector AURIGA during time spans which include the arrival time of a GRB
burst. The AURIGA detector \[\text{[13]}\] is an Al5056 resonant bar of about 2.3 tons with a typical noise temperature of 7 mK and a bandwidth of about 1 Hz. The detector is sensitive to gravitational waves (g.w.) signals over 1 Hz bandwidth around each one of its two resonant frequencies i.e., 913 and 931 Hz. The sensitivity of a g.w. detector can be conveniently expressed by the quantity \(h_{\text{min}}\), representing the minimum g.w. amplitude detectable at SNR = 1, which for the AURIGA detector is \(h_{\text{min}} \sim 2.5 \times 10^{-19}\). The AURIGA sensitivity is enough to detect NS-NS, BH-NS and BH-BH mergers which take place within the Galaxy.

Previous work on GRB-GWB association appeared in the literature, specifically about a single GRB trigger \[\text{[13]}\], and about a set of GRB triggers \[\text{[4, 14]}\]. Here we present an experimental upper limit on such an association obtained with the AURIGA-BATSE data.

The plan of the paper is as follows: in Sec. II we discuss the two methods, coincidence and statistical searches, which we use for the association of the GRBs in the BATSE catalog with GWBs. The results obtained with the two methods are presented in Sec. III. In Sec. IV we analyze our results and discuss the possibilities opened by our methods for future searches.

II. SEARCH METHODS

The aim of our analysis is the search of an association between GRB and GWB arrival times within a time window \(W\). The analysis has been carried out by means of two different procedures which have been already discussed in the literature: \(i\) the correlation method \[\text{[17]}\] which is based on the coincidence between the arrival time of candidate gravitational events selected over a given threshold and the GRB trigger and \(ii\) a statistical method \[\text{[16]}\] which relies on a hypothesis test on a statistical variable representing the mean energy of the gravitational detector at the GRB trigger.

The coincidence window plays an important role in both the analysis. In fact the coincidence window should be wide enough to hold the delay distribution between the event lists used in this analysis belong to periods of satisfactory performance of AURIGA as described in Ref. \[\text{[22]}\]. It is worth noticing that the AURIGA event search checks each event against the expected signal template by means of a \(\chi^2\) test \[\text{[23]}\]. The event lists contain the information needed to describe a \(\delta\)-like signal namely, its time of arrival (in UTC units), the amplitude of the Fourier transform, and the detector noise level at that time. The BATSE data are taken from the 4B catalog by Meegan et al. available on Internet \[\text{[24]}\] and includes trigger time (in UTC units), right ascension and declination, error box and other information about every GRB triggered.

We label with \(t_A^{(i)}\) the estimate arrival time of the \(i\)-th candidate g.w. event detected by AURIGA and with \(t_k^{(k)}\) the trigger time of the \(k\)-th GRB detected by BATSE. A coincidence between the \(i\)-th AURIGA event and the \(k\)-th GRB trigger time is observed if \(|t_A^{(i)} - t_k^{(k)}| \leq W\), where \(W\) is the coincidence window.

The coincidences found have to be compared with the accidental coincidence background due to chance. Two standard methods to evaluate the probability of accidental events have been applied to the pair BATSE-AURIGA: \(i\) performing thousands time-shifts of the arrival time of one detector with respect to the other and looking for accidental coincidences at each shift \[\text{[23]}\]; \(ii\) assuming independent Poisson distribution of event times and using the mean measured rates to estimate the accidental rate \(n_a\)

\[n_a = N_A N_y \frac{\Delta T}{T},\]  

(1)
where \( N_A \) is the number of AURIGA events in the period \( T \), \( N_o \) the number of GRB events in the same period and \( \Delta T = 2W \) is the total amplitude of the coincidence window (see Appendix A for a proof of this basic relation). To avoid the problem of multiple coincidence within the same window we prefer to estimate the background with \( W = 1 \text{ s} \) to scale our results with the help of Eq. (1) to \( W = 5 \text{ s} \). The consistency of the two estimates ensures that the arrival times of the AURIGA-BATSE pair are distributed as Poisson random points even if the AURIGA event rate has been found to be not stationary [17,21].

B. Statistical search

The method of the statistical search has been proposed for a pair of interferometric detectors by L. S. Finn and coworkers [13]; here we slightly modify their approach to the case of a single resonant detector. The data we use are the AURIGA filtered data obtained by means of the Wiener filter matched to a \( \delta \)-like signal, without setting any threshold on their amplitudes. If a signal enters the detector at time \( t_0 \), its output can be written as

\[
y(t) = hf(t - t_0) + \eta(t),
\]

where \( f(t - t_0) \) is the normalized signal template of the Wiener filter, \( h \) its amplitude (i.e. \( f(0) = 1 \)), \( t_0 \) its arrival time and \( \eta(t) \) is a stochastic process with zero mean and correlation

\[
\langle \eta(t), \eta(t') \rangle = \sigma_h^2 f(t - t').
\]

Here \( \sigma_h^2 \) is the variance of the filter output in the absence of any signal and \( f(t) \) is a superposition of two exponentially damped oscillating functions [27] which can be approximately expressed as

\[
f(t) \approx e^{-|t|/\tau_w} \cos(\omega_0 t) \cos(\omega_B t) \quad (\omega_0 \tau_w \gg 1),
\]

where \( \tau_w \) is the Wiener filter decaying time (i.e. the inverse of the detector bandwidth), \( \omega_0 \) a center carrier frequency and \( \omega_B \) an amplitude modulation frequency. Typical values for the AURIGA detector are \( \tau_w \approx 1 \text{ s} \), \( \omega_0 \approx 920 \text{ Hz} \) and \( \omega_B \approx 20 \text{ Hz} \).

Let us define now the random variable \( X \), which represents an averaged measure of the energy released by a g.w. signal impinging on the detector at time \( t_0 \)

\[
X(t_0) = \frac{1}{2W} \int_{-W}^{W} \text{d}t \ |y(t - t_0)|^2,
\]

where \( t_0 \) is the center of the time window. If an association between GRB and gravitational waves exists, the filtered output of the gravitational waves detector, in periods just prior the GRB (“on-source” population) will differ (statistically) from the output at other times (“off-source” population).

A statistically significant difference between on- and off-source populations clearly supports a GWB-GRB association. For each of the \( N_{on} \) GRB trigger we compute \( X(t_1^{(k)}) \), \( k = 1, \ldots, N_{on} \) which forms the \( \chi_{on} \) set of on-source events, and construct a complementary set \( \chi_{off} \) with \( N_{off} \) off-source events using windows before and after the trigger. The sets \( \chi_{on} \) and \( \chi_{off} \) are samples drawn from the populations whose distributions we denote \( p_{on} \) and \( p_{off} \). The off-source events are taken in periods not correlated with the trigger time, at a distance in time greater than \( 10^3 \) seconds from GRB the trigger, both before and after it; this should be sufficient to have a fair sample of the off-source events as any GRB-GWB association is reasonably excluded.

For windows \( W \) greater than the Wiener filter characteristic time, the central limit theorem implies that \( p_{off} \) is a normal distribution.

Now suppose that the GWBs fall within the window \( W \) opened around the GRB trigger time and that the SNR of the gravitational signal associated with the GRB (averaged over the source population) is smaller than one, then \( p_{on} \) is also a normal distribution with mean

\[
\mu_{on} = \mu_{off} + E \left[ \frac{1}{2W} \int_{-W}^{W} \text{d}t |h f(t - t_\gamma)|^2 \right] 
\]

\[
\approx \mu_{off} + \left( \frac{\tau_w}{8W} \right) \text{E}[h^2] \quad (W \gg \tau_w),
\]

where \( \text{E}[\cdot] \) is the average over the astrophysical source population of GRBs.

The basic idea behind the statistical approach is a hypothesis testing where the null hypothesis \( H_0 \) to test is the equivalence of the off-source and on-source distributions:

\[
H_0 : \ p_{off}(X) = p_{on}(X).
\]

The rejection of \( H_0 \) clearly supports a GWB-GRB association. Since \( p_{on} \) and \( p_{off} \) are normal and could differ only in their mean values, we can test \( H_0 \) by the Student’s \( t \)-test [28].

The \( t \) statistic is defined from \( \chi_{on} \) and \( \chi_{off} \) by

\[
t = \frac{\hat{\mu}_{on} - \hat{\mu}_{off}}{\Sigma} \left( \frac{N_{on} N_{off}}{N_{on} + N_{off}} \right)^{1/2},
\]

\[
\Sigma^2 = \frac{(N_{on} - 1) \sigma_{on}^2 + (N_{off} - 1) \sigma_{off}^2}{N_{on} + N_{off} - 2},
\]

where \( \hat{\mu}_{on} \) and \( \hat{\mu}_{off} \) (\( \sigma_{on}^2 \) and \( \sigma_{off}^2 \)) are the sample means (variances) of \( \chi_{on} \) and \( \chi_{off} \), respectively.

The expected value of \( t \) averaged on the source population and on the filtered output of the detector is

\[
\mu_t = E[t] = \left( \frac{\tau_w}{8W} \right) \frac{\text{E}[h^2]}{\sigma} \left( \frac{N_{on} N_{off}}{N_{on} + N_{off}} \right)^{1/2},
\]

where \( \sigma = E[\Sigma] \).
Let us consider the upper limit on \( E[h^2] \) in the assumption \( \mathcal{H}_0 \) is true: in this case the most probable value of \( E[h^2] \) is zero. From Eq. \( 10 \) we get

\[
\frac{t_w}{8W} E[h^2] \leq \mu_{t,\text{max}} \frac{N_{\text{on}} + N_{\text{off}}}{N_{\text{on}} N_{\text{off}}} =
\]

\[
= \left\{ \begin{array}{ll}
\mu_{t,\text{max}} \sqrt{2/N_{\gamma}} & (N_{\text{on}} = N_{\text{off}} = N_{\gamma}) \\
\mu_{t,\text{max}}/\sqrt{N_{\text{on}}} & (N_{\text{off}} \gg N_{\text{on}})
\end{array} \right.
\]

where \( \mu_{t,\text{max}} \) is the upper limit. Assuming \( \mathcal{H}_0 \) true and \( N_{\text{off}} \gg N_{\text{on}} \), we have

\[
E[h^2] \leq h_{\text{max}}^2 = \frac{8W}{t_w} \frac{\mu_{t,\text{max}}}{\sqrt{N_{\text{on}}}} \sigma ;
\]

\[
(12)
\]

the value of \( \mu_{t,\text{max}} \) can be deduced by the selected confidence level. In this way we are able to set an upper limit on the average amplitude of gravitational signals associated with GRBs

\[
h_{\text{RMS}}^2 \leq (1.4 \times 10^{-18})^2 \frac{W}{5s} \frac{(t_w/1s)^{-1}}{1.96} \times \sqrt{N_{\text{on}}}^{-1/2} \sigma \left(\frac{N_{\text{on}}}{100}\right)^{-1/2} \frac{5 \times 10^{-19}}{2}.
\]

\[
(13)
\]

III. RESULTS

The AURIGA data used in the two analysis are relative to the years 1997 and 1998, and the number of GRBs which fall into the AURIGA data taking periods is 120.

A. Coincidence search

Within the window of 5 seconds we have found 2 events in coincidence. This experimental result has to be compared with the number of coincidences due to chance.

The shifts method consists of \( 10^4 \) time shifts of the coincidence window; for each one we compute the number of coincidences; if the candidate GWB arrival times and the GRB triggers can both be modeled as Poisson random points \( [27] \), the number of accidental coincidences is fitted to a Poisson curve: from the fit we get the expected number of coincidences due to chance. The results are summarized in Fig. \( 2 \). From the fit we obtain a value of mean expected accidental coincidences of \( n_a = 2.57 \pm 0.04 \).

Another approach to evaluate the number of accidental coincidences is given by Eq. \( 11 \) that holds in case of Poisson random points \( [27] \). The total number of AURIGA events and GRB triggers are respectively \( N_A = 26816, N_{\gamma} = 120 \), assuming \( \Delta T = 2W = 10 \) s and \( T = 1.32 \times 10^5 \) s we get \( n_a = 2.4 \pm 0.2 \). The error on \( n_a \) can be easily estimated assuming the Poisson statistics for the fluctuations on \( N_A \) and \( N_{\gamma} \) i.e. \( \sqrt{N_A} \) and \( \sqrt{N_{\gamma}} \) respectively. The two estimates of the accidental number of coincidences are in good agreement and demonstrate that the two event rates are uncorrelated. We conclude that the 2 coincidence found are due to chance.

B. Statistical search

We have first tested the method using a Monte Carlo simulation by adding \( N_{\text{on}} \) signals at fixed SNR over a gaussian noise generated by means of a noise model with the same parameters of the AURIGA detector. The simulated detector output is then fed to the same data filtering procedures used for the AURIGA experimental data. We have generated 120 \( \delta \)-like signals with SNR 3, 2 and 1 and 1000 with SNR 1 storing their true arrival times \( t_0 \). The signals have been superimposed to stationary gaussian noise and we have then formed the on and off populations and calculated the \( t \) value. The probability \( P(t) \) that the \( t \) value obtained is due to chance is reported in Table \( 1 \). For signal with \( SNR = 3 \) and \( SNR = 2 \) this probability is very small and the value of \( t \) is statistically significant, showing a GRB-GWB association. On the other hand, for signals with \( SNR = 1 \), we have to increase the number of GRBs to 1000 to get a statistical significance.

The \( \chi_{\text{off}} \) and \( \chi_{\text{on}} \) sets relative to the AURIGA-BATSE data between 1997 and 1998 are given in the histograms of Fig. \( 2 \) and Fig. \( 3 \) where the non gaussian tails are due to the non stationarity of the AURIGA noise. The value of the Student’s t-test obtained is \( t = 0.58 \), which corresponds to a probability of 0.28 that it is due to chance.

As we conclude that there is no evidence of an association of GRB-GWB, we can put a constraint on gravitational radiation emitted from GRBs averaged over the source population, \( h_{\text{RMS}} \). The value of \( \mu_{t,\text{max}} \) can be found setting a confidence level C.L. and solving the equation:

\[
\int_{-\infty}^{\mu_{t,\text{max}}} f_d(t) \, dt = \text{C.L.} ,
\]

\[
(14)
\]

where \( f_d(t) \) is the distribution function for the Student’s \( t \) with \( d \) degrees of freedom. Notice that for high \( d \), \( f_d(t) \) tends to a normal curve with zero mean and unitary variance. If we choose C.L.=95\% and set \( d = N_{\text{on}} + N_{\text{off}} - 2 \), we get \( \mu_{t,\text{max}} = 1.65 \). Using Eq. \( 13 \) we set the following upper limit on the averaged g.w. amplitude associated with GRBs:

\[
h_{\text{RMS}} \leq 1.5 \times 10^{-18}.
\]

\[
(15)
\]

IV. DISCUSSION

The two search methods reported in this paper show no evidence of a correlation between \( \gamma \)-ray bursts and gravitational waves: First, the two coincidences found with
the coincidence search have no statistical significance and can be reconsidered to chance. Next, the statistical method does not lead to a statistically significant value of the Student’s $t$. However we were able to put an upper limit on gravitational signals associated to the GRBs averaged over the source population, $h_{\text{RMS}} \leq 1.5 \times 10^{-18}$.

The existence of burst-like excitations (usually referred as non-gaussian noise) in the AURIGA data is well known, and here we deal with this noise selecting the periods of time when the detector was operating in a satisfactory way. The vetoing procedure of data dominated by non gaussian noise has been set by the AURIGA data analysis [28] and the resulting duty cycle is about 40% of the total operating time. Care was also taken to cope with non stationarities of the AURIGA noise that give rise to the non gaussian tails in Fig. 4 and Fig. 5 excluding such tails from the fits we used to estimate the statistical parameter $t$. However, it is important to notice that even with the above limitations the analysis has reached a level of sensitivity which is astrophysically of some interest. In order to get better upper limits we have also explored the possibility of using in our analysis the incoming direction of GRBs but we have found no significant improvements (see Appendix B).

To increase the confidence in our results we have applied the non-parametric Mann-Whitney $u$-test [29] to the sample sets of the statistical search, obtaining a second confirmation of the null hypothesis. Ranking the events $\chi_{\text{off}} \oplus \chi_{\text{on}}$ in increasing order, we get a statistical parameter $z = 1.59$, smaller than the critical value $z = 1.95$ imposed by the fixed C.L. of 95%.

The upper limit we have set can be improved in the future simply by the increasing of common data taking periods of AURIGA ($N_{\text{on}}$ increases) and the new experiment HETE-II (High Energy Transient Explorer [30]) which is going to substitute by now the wasted BATSE satellite in the GRB search. Another possibility to reach more astrophysically interesting sensitivities is the upgrade of the AURIGA detector which is now in progress [31].

The predicted sensitivity and bandwidth of the AURIGA detector equipped with the new read out system would be respectively $h_{\text{min}} \approx 8 \times 10^{-20}$ and $t_w^{-1} \approx 10$ Hz. This sensitivity, together with an enhancement of noise stationarity and duty cycle of the detector, would correspond to the lowering of the upper limit $h_{\text{RMS}}$ of about 2 order of magnitude in one year of correlation analysis.

**APPENDIX A:**

To derive Eq. 1 let us consider $n$ points random distributed in a time interval $[0, T]$. The probability $P(\{k_a, t_a\})$ that $k_a$ points lies in the time window $t_a = t_2 - t_1$ is given by the binomial distribution [27] with probability $p = t_a/T$ that a single point lie in $t_a$. If $n \gg 1$ and $t_a \ll T$, using the Poisson theorem we get

$$P(\{k_a, t_a\}) = e^{-nt_a/T} \frac{(nt_a/T)^{k_a}}{k_a!}.$$  

(A1)

For $m$ not overlapping windows, it can be demonstrated that in the limit of $n \to \infty$, $T \to \infty$ and $n/T$ constant, the probability of $\{k_1 \text{ points in } t_1\}, \ldots, \{k_m \text{ points in } t_m\}$ is [27]

$$P(\{k_1, t_1\}, \ldots,\{k_m, t_m\}) = \prod_{i=1}^{m} e^{-nt_i/T} \frac{(nt_i/T)^{k_i}}{k_i!} = \prod_{i=1}^{m} P(\{k_i, t_i\})$$  

(A2)

showing that the events $\{k_1 \text{ points in } t_1\}$ and $\{k_j \text{ points in } t_j\}$ are independent for every $i$ and $j$. Substituting $m = N, t_i = 2W = \Delta T$ and $n = N_A$ we get

$$P(\{k_1, \Delta T\}, \ldots, \{k_N, \Delta T\}) =$$

$$= e^{-N, N_A \Delta T/T} \left( \frac{N_A \Delta T}{T} \right)^{k} \prod_{i=1}^{N} \frac{1}{k_i!}.$$  

(A3)

We can now evaluate the probability $P(k)$ to have $k$ coincidences, that can be obtained taking into account, at fixed $k$, all the possible sets $\{k_i\} = \{k_1, \cdots, k_N\}$ with the constrain $\sum_{i=1}^{N} k_i = k$; we get

$$P(k) = \sum_{\{k_i\}} P(\{k_1, \Delta T\}, \ldots, \{k_N, \Delta T\})$$

$$= e^{-N_A \Delta T/T} \left( \frac{N_A \Delta T}{T} \right)^{k} \prod_{i=1}^{N} \frac{1}{k_i!}$$

$$= e^{-N_A \Delta T/T} \left( \frac{N_A \Delta T}{T} \right)^{k} \frac{N!}{k!}.$$  

(A4)

Rearranging the factors in Eq. A4 we get a Poisson distribution as in Eq. 4. The last equality can be easily demonstrated using the multinomial expansion relation [22]

$$\left( \sum_{i=1}^{N} x_i \right)^k = k! \prod_{i=1}^{N} \frac{x_i^{k_i}}{k_i!}.$$  

(A5)

and substituting $x_i = 1, i = 1 \cdots N$.

**APPENDIX B:**

As the incoming direction of GRBs is known, one may wonder if a selection of the GRBs based on a cutoff on the AURIGA antenna pattern, averaged on the unknown polarizations of the GWB, could increase the sensitivity of our analysis.

The antenna pattern of a resonant bar detector is given by
\[ F(\theta) = 1 - (\hat{k} \cdot \hat{z})^2 = \sin^2(\theta) \], \hspace{1cm} (B1)

where \( \hat{k} \) and \( \hat{z} \) are unitary vectors parallel to the GRB direction and the antenna bar axis and \( \theta \) is the angle between \( \hat{k} \) and \( \hat{z} \). Sources that fall outside the two cones, which are defined by the equation \( \hat{k} \cdot \hat{z} \geq \cos(\xi) \), have a figure pattern \( F(\theta) \geq \sin^2(\xi) \) and therefore the average energy associated with these g.w. sources is

\[ E^\xi[h^2] \propto \frac{15}{16} \int_{\Omega_\xi} F^2(\theta) \frac{d\Omega}{4\pi} \]

\[ = \frac{75}{64} \left[ \sin(\xi) + \frac{1}{6} \sin(3\xi) + \frac{1}{50} \sin(5\xi) \right], \hspace{1cm} (B2) \]

where the solid angle \( \Omega_\xi \) is defined by \( F(\theta) \geq F(\xi) \). Moreover, as GRBs are isotropically distributed over the sky, the cutoff on the figure pattern decreases the number of available GRB: \( N_{\xi_{\text{on}}}^\xi = N_{\text{on}} \sin(\xi) \). Therefore the net effect of this cutoff on the expected value of \( t \) in Eq. 10 is

\[ \mu^\xi_{\text{on}} = \mu_{\text{on}} \frac{75}{64} \left[ \sin(\xi) + \frac{1}{6} \sin(3\xi) + \frac{1}{50} \sin(5\xi) \right] \times \sin(\xi)^{1/2}. \hspace{1cm} (B3) \]

The function \( \mu^\xi_{\text{on}}/\mu_{\text{on}} \) is a continuously increasing function in the range \( \xi \in [0, \pi/2] \), and \( \mu^\xi_{\text{on}}/\mu_{\text{on}} = 1 \) at \( \xi = \pi/2 \) (i.e. the whole solid angle). We must conclude that a cutoff on the GRB direction does not enhance the sensitivity of the statistical search.

![FIG. 1. Poisson fit of the data obtained with the shifts method.](image1)

![FIG. 2. Gaussian fit of the off-source set of the statistical search.](image2)

TABLE I. Results of a Monte Carlo simulation of the statistical search. The column \( P(t) \) is the probability that the estimated \( t \) is due to chance.

| SNR of the generated signals | \( N_{\text{on}} \) | \( t \) | \( P(t) \) |
|-----------------------------|----------------|--------|---------|
| 3                           | 120            | 8.1    | \(< 10^{-9}\) |
| 2                           | 120            | 3.4    | \(4 \times 10^{-4}\) |
| 1                           | 120            | 0.4    | \(3 \times 10^{-1}\) |
| 1                           | 1000           | 3.6    | \(10^{-4}\) |

TABLE II. Results of the fits on the \( \chi_{\text{off}} \) and \( \chi_{\text{on}} \) sets with a gaussian distribution.

| \( N \) | \( \mu \) | \( \sigma \) |
|--------|----------|----------|
| off    | 2206     | \((0.97 \pm 0.02) \times 10^{-36}\) | \((0.37 \pm 0.02) \times 10^{-36}\) |
| on     | 120      | \((0.99 \pm 0.06) \times 10^{-36}\) | \((0.33 \pm 0.06) \times 10^{-36}\) |

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FIG. 3. Gaussian fit of the on-source set of the statistical search.

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