A KEPLERIAN-LIKE DISK AROUND THE FORMING O-TYPE STAR AFGL 4176

KATHARINE G. JOHNSTON1,2, THOMAS P. ROBITAILLE2, HENRIK BEUTHER2, HENDRIK LINZ2, PAUL BOLEY3, ROLF KUIPER4,5, ERIC KETO5, MELVIN G. HOARE1, and ROY VAN BOEKKEL2

1 School of Physics & Astronomy, E.C. Stoner Building, The University of Leeds, Leeds LS2 9JT, UK; k.g.johnston@leeds.ac.uk
2 Institute of Astronomy and Astrophysics, Eberhard Karls University, Tübingen, Auf der Morgenstelle 10, D-72076 Tübingen, Germany
3 Max Planck Institute for Astronomy, Königstuhl 17, D-69117 Heidelberg, Germany
4 Harvard-Smithsonian Center for Astrophysics, 60 Garden Street, Cambridge, MA 02138, USA
5 Ural Federal University, Astronomical Observatory, 51 pr. Lenina, Ekaterinburg, Russia

Received 2015 September 11; accepted 2015 September 28; published 2015 October 29

ABSTRACT

We present Atacama Large Millimeter/submillimeter Array line and continuum observations at 1.2 mm with ∼0′′3 resolution that uncover a Keplerian-like disk around the forming O-type star AFGL 4176. The continuum emission from the disk at 1.21 mm (source mm1) has a deconvolved size of 870 ± 110 AU × 330 ± 300 AU and arises from a structure ∼8 $M_\odot$ in mass, calculated assuming a dust temperature of 190 K. The first-moment maps, pixel-to-pixel line modeling, assuming local thermodynamic equilibrium (LTE), and position–velocity diagrams of the CH$_3$CN $J = 13–12$ K-line emission all show a velocity gradient along the major axis of the source, coupled with an increase in velocity at small radii, consistent with Keplerian-like rotation. The LTE line modeling shows that where CH$_3$CN $J = 13–12$ is excited, the temperatures in the disk range from ∼70 to at least 300 K and that the H$_2$ column density peaks at $2.8 \times 10^{22}$ cm$^{-2}$. In addition, we present Atacama Pathfinder Experiment $^{12}$CO observations that show a large-scale outflow from AFGL 4176 perpendicular to the major axis of mm1, supporting the disk interpretation. Finally, we present a radiative transfer model of a Keplerian disk surrounding an O7 star, with a disk mass and radius of 12 $M_\odot$ and 2000 AU that reproduces the line and continuum data, further supporting our conclusion that our observations have uncovered a Keplerian-like disk around an O-type star.

Key words: circumstellar matter – ISM: jets and outflows – radiative transfer – stars: formation – stars: massive – techniques: interferometric

1. INTRODUCTION

The process of star formation is often pictured as a young star being fed by a disk formed as a result of angular momentum conservation over a period of tens of thousands of millions of years. Yet, this picture has been mainly derived from observations of disks that have practically finished accreting. In reality, it has been confirmed only relatively recently that stable, rotationally supported, Keplerian disks around low-mass Class 0/I protostars exist at earlier times, when the star and its disk are concealed deeply in an accreting envelope (e.g., Brinch et al. 2007; Jørgensen et al. 2009; Tobin et al. 2012; Murillo et al. 2013).

The situation is even less clear when it comes to forming massive stars, which go on to reach a final mass of 8 $M_\odot$ and above. As they form so quickly (t$_{\text{form}}$ ∼ 10$^4$–10$^5$ years), they spend the entirety of their formation still embedded in a surrounding envelope. Are these protostars the high-mass equivalents of low-mass Class 0/I objects, and do they also have disks? Many of the energetic feedback mechanisms associated with massive stars, such as radiation pressure and at later times stellar winds and photoionization, would halt the vast amount of accretion required to form these stars. In theory, these can be bypassed by the existence of a disk (e.g., Yorke & Sonnhalter 2002; Krumholz et al. 2009; Kuiper et al. 2010, 2011), allowing the radiation pressure, winds, or hot ionized gas to be channeled away along the axis perpendicular to the disk, where densities are lower.

Recent observations of early B-type (proto)stars have responded to this prediction with detections of disk candidates that have kinematics that appear to be dominated by the central protostar and are stable (e.g., Johnston et al. 2011; Sánchez-Monge et al. 2013; Beltrán et al. 2014; Cesaroni et al. 2014). Yet, the number of disks discovered so far only constitutes a handful (Cesaroni et al. 2007; Beltrán et al. 2011), as many of the observed rotating structures, instead referred to as toroids, are too large and rotate too slowly to be in centrifugal equilibrium. Stepping up to O-type stars, there are even fewer candidates (e.g., Wang et al. 2012; Hunter et al. 2014; Zapata et al. 2015); in these cases, the detections are based on a velocity gradient across the source and an outflow or jet that is projected perpendicular to the candidate disk plane.

In this Letter, we present Atacama Large Millimeter/submillimeter Array (ALMA) observations that trace a Keplerian-like disk toward the infrared source AFGL 4176, which constitutes the best observational example of an O-type protostar with a disk to date.

AFGL 4176 (G308.918+0.123, IRAS 13395–6153) is a forming massive star with coordinates 13h43m01s.22 +35°34′31″ (FK5 J2000) embedded in a star-forming region with a total luminosity of ∼10$^5$ $L_\odot$ (d = 4.2 kpc; Green & McClure-Griffiths 2011; Boley et al. 2012). The source lies at the northern edge of an H II region (Caswell et al. 1992; Ellingsen et al. 2005; Shabala et al. 2006) that peaks ∼4″ to the south of AFGL 4176, likely powered by another star of spectral type O9. Four 6.7 GHz Class II methanol maser spots with an extent of 840 AU at 4.2 kpc lie in close proximity to AFGL 4176 along a line with position angle (PA) ∼ −35° (Phillips et al. 1998). NH$_3$ observations have uncovered that the star is embedded in a large-scale rotating toroid with a radius of ∼0.7 pc (Johnston et al. 2014). The large-scale continuum emission at 1.2 mm (Beltrán et al. 2006) traces a dense core of 0.8 pc and 890 $M_\odot$ at 4.2 kpc, and knots of shocked H$_2$ emission have been detected (De Buizer et al. 2009).
suggested the presence of an outflow. Boley et al. (2012) have modeled the spectral energy distribution and mid-IR interferometric observations of AFGL 4176, finding that the latter required a non-spherically symmetric model to adequately fit the data. They interpreted the mid-IR visibilities as a combination of a disk-like structure with radius of 660 AU and a spherically symmetric Gaussian halo with an FWHM of ~600 AU (Boley et al. 2012, 2013). Finally, Ilee et al. (2013) found that the 2.3 μm CO bandhead emission toward AFGL 4176 is consistent with the inner ~10 AU of a Keplerian disk.

2. OBSERVATIONS

2.1. ALMA Observations

We observed AFGL 4176 with the 12 m antenna array of ALMA during Cycle 1, under program 2012.1.00469.S (PI: Johnston). The observations were carried out on 2014 August 16 and 17 in dual-polarization mode in Band 6 (~250 GHz or 1.2 mm) under good weather conditions (precipitable water vapor, PWV ~ 1.32 and 1.14 mm, respectively). AFGL 4176 was observed with one pointing centered on 13°43′01″08′′−62°08′55″5′′ (FK5 J2000). Two wide and narrow spectral windows (spws) were observed with respective widths of 1.875 GHz and 468.750 MHz. The two wide spws were centered at 254.043 GHz, while the two narrow spws were centered at frequencies of 239.072 and 256.349 GHz. The spectral resolution was 1129 kHz (1.41 and 1.33 km s⁻¹) and 282 kHz (0.354 and 0.330 km s⁻¹), respectively. Thirty-nine antennas were included in the array, of which 36 had useful data. Baseline lengths were 14.4–1210.8 m, providing a largest angular scale of ~18″. The primary beam size was 22″7–24″4. The bandpass calibrators were J1617-5848 and J1427-4206 and the absolute flux calibrators were Titan and Ceres on August 16 and 17, respectively. Phase/gain calibrators were J1308-6707 and J1329-5608 for both days. The flux calibration uncertainty was estimated to be ≤20%.

Calibration was carried out using the Common Astronomy Software Applications (CASA) version 4.2.1 via the delivered pipeline script. We improved the resulting images by self-calibration of the continuum and applied these solutions to the line data. The continuum images were made using 1.4 GHz bandwidth of line-free channels across all spws. The central frequency of the combined continuum emission is 247.689 GHz (1.210 mm). Imaging was carried out using Briggs weighting with a robust parameter of 0.5. The noise in the continuum image is 78 μJy beam⁻¹ in a beam of 0″28 × 0″24, PA = −30°2. In this Letter, we also present the detected K-ladder transitions of CH₃CN J = 13–12. The K = 0–8 transitions were observed within the 239.072 GHz band. For the K = 2–8 images presented below, the noise ranges between 3.4 and 6.2 mJy beam⁻¹ in a beam of 0″30 × 0″28, PA = 37°7–37°9.

2.2. Atacama Pathfinder Experiment (APEX) Observations

We observed AFGL 4176 with the APEX⁶ 12 m antenna for program M0020_89 during the night of 2012 April 20–21 under very good weather (PWV ~ 0.4 mm). The Swedish Heterodyne Facility Instrument APEX-2 receiver was tuned to 12CO(3–2) at 345.79599 GHz in the lower sideband, providing a beam size of 18″. On-the-fly maps with an extent of 2.5 × 2.5′ were taken in two perpendicular scan directions to reduce scanning artifacts. The XFFTS2 backend provided a channel separation of 0.1984 km s⁻¹. The achieved rms in the spectra in the central part of the map is around 0.12 K. The data reduction was performed within GILDAS/CLASS.⁷ The units of the spectral cube are in corrected antenna temperatures (T_A^*).

3. RESULTS AND DISCUSSION

In Figure 1, we present the 1.21 mm continuum emission from the AFGL 4176 region. The emission is dominated by mm1, with a peak position in the non-self-calibrated image of 13°43′01″093−62°08′51″25 (FK5 J2000), coincident with the position of AFGL 4176 in 2MASS (with ~0″1 positional uncertainty). Performing a Gaussian fit to mm1, we determined a peak flux of 37 ± 2 mJy beam⁻¹, an integrated flux of 50 ± 4 mJy, a deconvolved size of 0″21 ± 0″03 × 0″08 ± 0″07 (870 ± 110 AU × 330 ± 300 AU), and PA = 59° ± 17°. Using the equations of Hildebrand (1983), as well as assuming 190 K (the average temperature of the disk derived from the CASSIS LTE line modeling below), a gas-to-dust ratio of 154 (Draine 2011), and a dust opacity at 1.21 mm of 0.24 cm² g⁻¹ (Draine 2003a, 2003b, with RV = 5.5), we derive a gas mass and peak column density of 8 M⊙ and 8 × 10²⁴ cm⁻² for mm1. As well as a compact component, mm1 includes low-lying 5σ–10σ emission extending NW, perpendicular to the major axis of mm1, that joins it to mm2. The remaining millimeter sources lie in the NW and SE quadrants around mm1, including two compact sources, mm2 and mm4, on opposite sides of mm1. Full observed properties of all detected 1.21 mm sources will be given in a upcoming paper (K. G. Johnston et al. 2016, in preparation).

Figure 2 shows the integrated 12CO J = 3–2 emission from the red- and blueshifted high-velocity wings of the bipolar outflow from AFGL 4176 observed by APEX, overlaid upon a 12CO zero-moment map shown in grayscale. The outflow lobes are oriented roughly NW–SE, perpendicular to the PA of the source mm1, shown as a dashed line in Figure 2. This relative geometry suggests mm1 is a disk driving the large-scale outflow seen in 12CO.

The kinematics of the gas in mm1 traced by first- and second-moment maps (intensity-weighted velocity and line-width fields, respectively) of CH₃CN J = 13–12 K = 3 emission are shown in panels (a) and (b) of Figure 3. The velocity field of the K = 3 line in panel (a) shows a clear velocity gradient along the major axis of mm1, which is also present in lines K = 2–8 and on arcminute scales in NH₃ (Johnston et al. 2014). The velocity field shown in Figure 3 is similar to that expected from disks in near-Keplerian rotation (e.g., HD 100546 and TW Hya, respectively; Hughes et al. 2011; Pineda et al. 2014). For instance, there is a quick change from blue- to redshifted emission when crossing the minor axis of the source. In addition, the emission in the most blue- and redshifted channels is found close to the continuum peak position, whereas the lower-velocity gas is more extended. The line widths across mm1 in panel (b) peak toward the continuum

---

⁶ APEX is a collaboration between the Max-Planck-Institut für Radioastronomie, the European Southern Observatory, and the Onsala Space Observatory.

⁷ http://www.iram.fr/IRAMFR/GILDAS

http://www.ipac.caltech.edu/2mass/releases/allsky/doc/sec1_6b.html#parsecpsour
peak, consistent with beam-averaging of high-velocity red- and blueshifted emission. There is extended CH₃CN emission to the NW of mm1 with near-systemic velocities \( \left( v_{\text{LSR}} \sim -52 \text{ km s}^{-1} \right) \), as well as blueshifted gas associated with mm2, which would be consistent with this source being associated with the blueshifted lobe seen in \(^{12}\)CO.

Figures 3(c)–(f) present the results of modeling the CH₃CN and CH₃C≡N lines \( J = 13–12 \) \( K \)-ladders in the spectrum associated with each pixel, making a map of excitation temperature \( T_{\text{ex}} \), velocity \( v_{\text{LSR}} \), linewidth \( \Delta v_{\text{FWHM}} \), and column density \( N_{\text{mol}} \) across the source. The modeling was carried out using CASSIS⁹ and the JPL molecular spectroscopy database¹⁰ using the Markov Chain Monte Carlo \( \chi^2 \) minimization method and assuming local thermodynamic equilibrium (LTE). Given \( N_{\text{mol}} \) and \( T_{\text{ex}} \), CASSIS determines an optical depth that is included in each model. We assume a \([^{13}\text{C}/^{12}\text{C}] \) abundance ratio of 60 (assuming the distance to the Galactic center is 8.4 kpc and thus a Galactocentric distance to AFGL 4176 of 6.6 kpc; Milam et al. 2005; Reid et al. 2009), a source size of 0″.53 from the fitted size of the \( K = 0 \), 1 zero-moment map, and initial parameter ranges \( T_{\text{ex}} = 50–350 \text{ K} \), \( v_{\text{LSR}} = -57 \) to \( -47 \text{ km s}^{-1} \), \( \Delta v_{\text{FWHM}} = 0.5–10 \text{ km s}^{-1} \), and \( N_{\text{mol}} = 1 \times 10^{14}–1 \times 10^{17} \text{ cm}^{-2} \). The resultant temperature shown in panel (c) peaks at the continuum peak position and has values ranging between 74 and 294 K. There is also a secondary temperature peak toward mm2. Panels (d) and (e) show similar results to the first- and second-moment maps shown in panels (a) and (b); however, an increase in linewidth can be seen more clearly toward mm2. The column density of CH₃CN peaks at \( 2.8 \times 10^{17} \text{ cm}^{-2} \) toward the center of mm1, with a slight elongation along the major axis. Using the assumed size above and an abundance of CH₃CN of \( 10^{-8} \) (e.g., Gibb et al. 2000; Bisschop et al. 2007), this corresponds to \( H_2 \) column and volume densities of \( 2.8 \times 10^{24} \text{ cm}^{-2} \) and \( >8 \times 10^7 \text{ cm}^{-3} \), respectively.

Having established that a rotating structure, likely a disk, surrounds AFGL 4176 mm1, we can compare it with the geometry of the circumstellar material determined from modeling observations from the MID-infrared interferometric Instrument (MIDI; Boley et al. 2012, 2013). The PA of mm1 (\( \sim 60^\circ \)) is not aligned with the \( 112^\circ \) determined from the MIDI visibilities. Although these PA are not exactly orthogonal, these two observations may be reconciled if the MIDI observations are instead tracing heated dust in the outflow cavity, such as found for W33 A (de Wit et al. 2010). We also note that although offset from the continuum peak of mm1, the four Class II methanol maser spots detected by Phillips et al. (1998), shown in Figure 3(f), lie along a line almost perpendicular to the disk.

To investigate the dynamics of mm1 further, we present position–velocity (PV) diagrams of the CH₃C≡N lines \( J = 13–12 \) \( K \)-lines that had good signal-to-noise and were not contaminated by envelope emission. The PV diagrams of the \( K = 2 \), 4, and 6 lines exhibit the “butterfly” shape expected from a Keplerian-like rotation curve that increases in velocity close to the central object and is more extended close to the systemic velocity. It bears close resemblance to the PV diagrams observed toward the forming early B-type star IRAS 20126+4014, which has a disk in Keplerian-like rotation (Cesaroni et al. 2005, 2014), as well as the forming O-type stars NGC 6334 I(N) SMA 1b and IRAS 16547–4247 (Hunter et al. 2014 and Zapata et al. 2015, respectively). The velocity gradient of the \( K \)-lines becomes steeper at higher values of \( K \), thus the more-excited lines that trace hotter gas closer to the central object also trace higher velocities, which is expected from Keplerian-like rotation.

Interestingly, the PV diagrams of the lower \( K \)-lines shown in Figure 4 are asymmetrical about the systemic velocity, with the blueshifted emission being brighter. This is not observed

---

⁹ CASSIS is developed by IRAP-UPS/CNRS (http://cassis.irap.omp.eu).
¹⁰ http://spec.jpl.nasa.gov
however in the $K = 7, 8$ lines. A similar blue asymmetry is seen in the PV diagrams of the low-excitation $K$-lines of CH$_3$CN $J = 12$--$11$ observed by Cesaroni et al. (2014) and Hunter et al. (2014), whereas the high-excitation lines are either symmetrical or exhibit a red asymmetry. The CH$_3$CN $J = 19$--$18$ $K = 2$ PV diagram of G35.03+0.35 HMC A also shows this blue asymmetry (Beltrán et al. 2014). This has been suggested to be due to an asymmetric disk structure; however, the fact that so many objects exhibit the same features in their PV diagrams suggests this is more likely due to a radiative transfer effect and/or a geometry that is present in all sources.

To check whether a Keplerian-disk model is consistent with our observations, we ran a grid of self-consistent gas and dust radiative transfer models, where the line and continuum radiative transfer were performed using the codes MOLLIE (assuming LTE; Keto & Caselli 2010) and HYPERION (Robitaille 2011), using the Milky Way dust properties from Draine (2003a, 2003b) with $R_V = 5.5$, respectively. To fit the models to the line and continuum observations, we fit the profiles of the continuum and CH$_3$CN $J = 13$--$12$ $K = 2$--$8$ emission collapsed along the major and minor axes, as well as the integrated spectra for the lines. The models were convolved to the observed beam before fitting.

Panels (e)--(h) of Figure 4 show the CH$_3$CN $J = 13$--$12$ PV diagrams for a model that provides a good fit to the line and continuum data. The model consists of a Keplerian flared disk of radius 2000 AU (the inner radius is set to be the dust sublimation radius, which in this case is 31.3 AU), total gas mass $12M_\odot$, with a surface density decreasing as $r^{-1.5}$, and an inclination of $30^\circ$. The scale height of the disk is given by $z = 6.7(\varpi/100 \text{ AU})^{0.29} \text{ AU}$, where $\varpi$ is the cylindrical radius and the constants were determined in order for the disk to be in hydrostatic equilibrium. The model includes a rotationally flattened infalling envelope (Ulrich 1976) with an infall rate of $4.6 \times 10^{-4} M_\odot \text{ yr}^{-1}$ and an outer radius of 150,000 AU. The central object is a 25 $M_\odot$ zero-age main-sequence O7 star, with properties determined from Meynet & Maeder (2000) and Martins et al. (2005), specifically $T = 36,872$ K and $R = 6.71 R_\odot$. The Keplerian velocity field in the model accounts for the circumstellar mass interior to a given radius. The abundance of CH$_3$CN relative to H$_2$ is taken to be $10^{-4}$ at $>100$ K, dropping to $5 \times 10^{-9}$ for $90$ K $< T < 100$ K, and again to $10^{-10}$ for $<90$ K (Collings et al. 2004; Gerner et al. 2014). The grid of models as well as more detailed results will be presented in a future publication (K. G. Johnston et al. 2016, in preparation).

We note that the model disk mass of $12M_\odot$ differs from that derived using the single temperature of 190 K ($8M_\odot$) because it accounts for the expected variance in temperature and optical depth over the disk and is the best fit from a discrete set of disk masses. In addition, a Gaussian fit to the continuum model image gives a major FWHM of 880 AU, in good agreement with the fit to the observations (FWHM = 870 AU), which indicates the Gaussian fit underestimates the actual disk size.

4. CONCLUSIONS

We present ALMA observations that uncover the presence of a Keplerian-like disk around an O-type star, detected in both 1.21 mm continuum and CH$_3$CN. The velocity structure as traced
by CH$_3$CN moment maps and LTE line modeling shows a clear velocity gradient and evidence of higher velocities at small radii, as expected for a Keplerian disk. In addition, we present a radiative transfer model that agrees with our ALMA observations as expected for a Keplerian disk. In addition, we present a radiative transfer model that agrees with our ALMA observations as expected for a Keplerian disk. In addition, we present a radiative transfer model that agrees with our ALMA observations as expected for a Keplerian disk. In addition, we present a radiative transfer model that agrees with our ALMA observations as expected for a Keplerian disk. In addition, we present a radiative transfer model that agrees with our ALMA observations as expected for a Keplerian disk.

We thank the referee for insightful comments that helped improve this Letter. We are grateful to have been able to observe with ALMA and APEX on Llano de Chajnantor, Chile. ALMA is a partnership of ESO (representing its member states), NSF (USA) and NINS (Japan), together with NRC (Canada), NSC and ASIAA (Taiwan), and KASI (Republic of Korea), in cooperation with the Republic of Chile. The Joint ALMA Observatory is operated by ESO, AUI/NRAO and NAOJ. We thank Frank Wyrowski, Miguel Angel Requena Torres, and our ALMA contact scientist Edwige Chapillon. P.B. acknowledges support from the Russian Science Foundation, grant No. 15-12-10017.

REFERENCES

Beltrán, M. T., Brand, J., Cesaroni, R., et al. 2006, A&A, 447, 221
Beltrán, M. T., Cesaroni, R., Neri, R., & Codella, C. 2011, A&A, 525, A151
Beltrán, M. T., Sánchez-Monge, A., Cesaroni, R., et al. 2014, A&A, 571, A52
Bisschop, S. E., Jørgensen, J. K., van Dishoeck, E. F., & de Wachter, E. B. M. 2007, A&A, 465, 913
Boley, P. A., Linz, H., van Boekel, R., et al. 2012, A&A, 547, A88
Boley, P. A., Linz, H., van Boekel, R., et al. 2013, A&A, 557, C1
Brinch, C., Crapsi, A., Jørgensen, J. K., Hogerheijde, M. R., & Hill, T. 2007, A&A, 475, 915
Caselli, P., Kesteven, M. J., Stewart, R. T., Milne, D. K., & Haynes, R. F. 1992, ApJL, 399, L151
Cesaroni, R., Galli, D., Lodato, G., Walmsley, C. M., & Zhang, Q. 2007, in Protostars and Planets V, ed. B. Reipurth, D. Jewitt & K. Keil (Tucson, AZ: Univ. Arizona Press), 197
Cesaroni, R., Galli, D., Neri, R., & Walmsley, C. M. 2014, A&A, 566, A73
Cesaroni, R., Neri, R., Olmi, L., et al. 2005, A&A, 434, 1019
Collings, M. P., Anderson, M. A., Chen, R., et al. 2004, MNRAS, 354, 1133
De Buizer, J. M., Redman, R. O., Longmore, S. N., Caswell, J., & Feldman, P. A. 2009, A&A, 493, 127
de Wit, W. J., Hoare, M. G., Oudmaijer, R. D., & Lumsden, S. L. 2010, A&A, 515, A45
Draine, B. T. 2003a, ARA&A, 41, 241
Draine, B. T. 2003b, ApJ, 598, 1017
Draine, B. T. 2011, Physics of the Interstellar and Intergalactic Medium (Princeton, NJ: Princeton Univ. Press)
Ellingsen, S. P., Shabala, S. S., & Kurtz, S. E. 2005, MNRAS, 357, 1003
Gerner, T., Beuther, H., Semenov, D., et al. 2014, A&A, 563, A97
Gibb, E., Nummelin, A., Irvine, W. M., Whittet, D. C. B., & Bergman, P. 2000, ApJ, 545, 309
Green, J. A., & McClure-Griffiths, N. M. 2011, MNRAS, 417, 2500
Hildebrand, R. H. 1983, QJRAS, 24, 267
Hughes, A. M., Wilner, D. J., Andrews, S. M., Qi, C., & Hogerheijde, M. R. 2011, ApJ, 727, 85
Hunter, T. R., Brogan, C. L., Cyganowski, C. J., & Young, K. H. 2014, ApJ, 788, 187
Ilee, J. D., Wheelwright, H. E., Oudmaijer, R. D., et al. 2013, MNRAS, 429, 2960
Johnston, K. G., Beuther, H., Linz, H., et al. 2014, in The Labyrinth of Star Formation (Astrophysics and Space Science Proceedings, Vol. 36; Cham, Switzerland: Springer International), 413
Johnston, K. G., Keto, E., Robitaille, T. P., & Wood, K. 2011, MNRAS, 415, 2953
Jørgensen, J. K., van Dishoeck, E. F., Visser, R., et al. 2009, A&A, 507, 861
Keto, E., & Caselli, P. 2010, MNRAS, 402, 1625
Krumholz, M. R., Klein, R. I., McKee, C. F., Offner, S. S. R., & Cunningham, A. J. 2009, Sci, 323, 754
Kuiper, R., Klahr, H., Beuther, H., & Henning, T. 2010, ApJ, 722, 1556
Kuiper, R., Klahr, H., Beuther, H., & Henning, T. 2011, ApJ, 732, 20
Martins, F., Schaerer, D., & Hillier, D. J. 2005, A&A, 436, 1049
Meynet, G., & Maeder, A. 2000, A&A, 361, 101
Milam, S. N., Savage, C., Brewster, M. A., Zavits, L. M., & Wyckoff, S. 2005, ApJ, 634, 1126
Murillo, N. M., Lai, S.-P., Bruderer, S., Harsono, D., & van Dishoeck, E. F. 2013, A&A, 560, A103
Phillips, C. J., Norris, R. P., Ellingsen, S. P., & McCulloch, P. M. 1998, MNRAS, 300, 1131
Pineda, J. E., Quanz, S. P., Meru, F., et al. 2014, ApJL, 788, L34
Reid, M. J., Menten, K. M., Zheng, X. W., et al. 2009, ApJ, 700, 137
Robitaille, T. P. 2011, A&A, 536, A79
Sánchez-Monge, Á., Cesaroni, R., Beltrán, M. T., et al. 2013, A&A, 552, L10
Shabala, S. M., Ellingsen, S. P., Kurtz, S. E., & Forbes, L. K. 2006, MNRAS, 372, 457
Tobin, J. J., Hartmann, L., Chiang, H.-F., et al. 2012, Natur, 492, 83
Ulrich, R. K. 1976, ApJ, 210, 377
Wang, K.-S., van der Tak, F. F. S., & Hogerheijde, M. R. 2012, A&A, 543, A22
Yorke, H. W., & Sonnhalter, C. 2002, ApJ, 569, 846
Zapata, L. A., Palau, A., Galván-Madrid, R., et al. 2015, MNRAS, 447, 1826