Spatially Resolved Spitzer-IRS Spectral Maps of the Superwind in M82

P. Beirão1, L. Armus2, M. D. Lehnert3, P. Guillard3,4, T. Heckman5, B. Draine6, D. Hollenbach7, F. Walter8, K. Sheth9, J. D. Smith10, P. Shopbell11, F. Boulanger12, J. Surace2, C. Hoopes5, and C. Engelbracht13

1 Observatoire de Paris, LERMA, CNRS, 61 Av. de l’Observatoire, 75014 Paris, France
2 Spitzer Science Center, California Institute of Technology, MC 220-06, Pasadena, CA 91125
3 CNRS, UMR 7095, Institut d’Astrophysique de Paris, 98 bis boulevard Arago, 75014 Paris, France
4 Sorbonne Universités, UPMC Université Paris VI, 4 place Jussieu, 75005 Paris, France
5 Center for Astrophysical Sciences, Department of Physics and Astronomy, Johns Hopkins University, Baltimore, MD 21218, USA
6 Princeton University Observatory, Peyton Hall, Princeton, NJ 08544-1001, USA
7 SETI Institute, Mountain View, CA 94043, USA
8 Max-Planck Institut für Astronomie, Königstuhl 17, D-69117 Heidelberg, Germany
9 North American ALMA Science Center, National Radio Astronomy Observatory, 520 Edgemont Road, Charlottesville, VA 22901, USA
10 Ritter Astrophysical Observatory, University of Toledo, Toledo, OH 43606, USA
11 Astronomy Department, California Institute of Technology, MC 249-17, Pasadena, CA 91125
12 Institut d’Astrophysique Spatiale (IAS), UMR 8617, CNRS & Université Paris-Sud 11, Bâtiment 121, 91405, Orsay Cedex, France
13 Steward Observatory, University of Arizona, 933 North Cherry Avenue, Tucson, AZ 85721, USA

ABSTRACT

We have mapped the superwind/halo region of the nearby starburst galaxy M82 in the mid-infrared with Spitzer – IRS. The spectral regions covered include the H2S(1) − S(3), [NeII], [NeIII] emission lines and PAH features. We estimate the total warm H2 mass and the kinetic energy of the outflowing warm molecular gas to be between \(M_{\text{warm}} \approx 5 - 17 \times 10^6 M_{\odot}\) and \(E_K \approx 6 - 20 \times 10^{53}\) erg. Using the ratios of the 6.2, 7.7 and 11.3 micron PAH features in the IRS spectra, we are able to estimate the average size and ionization state of the small grains in the superwind. There are large variations in the PAH flux ratios throughout the wind region. The Northern part of the wind has a significant population of PAH’s with smaller 6.2/7.7 ratios than either the starburst disk or the Southern wind, indicating that on average, PAH emitters are larger and more ionized. The warm molecular gas to PAH flux ratios (H2/PAH) are enhanced in the outflow by factors of 10-100 as compared to the starburst disk. This enhancement in the H2/PAH ratio does not seem to follow the ionization of the atomic gas (as measured with the [NeIII]/[NeII] line flux ratio) in the outflow. This suggests that much of the warm H2 in the outflow is excited by shocks. The observed H2 line intensities can be reproduced with low velocity shocks (\(v < 40\) km s\(^{-1}\)) driven into moderately dense molecular gas (10\(^2 < n_H < 10^4\) cm\(^{-3}\)) entrained in the outflow.

Key words: galaxies: starburst — galaxies: individual (M82) — infrared: galaxies

1 INTRODUCTION

Galactic outflows, “superwinds”, are driven by hot gas and momentum generated by the combined mechanical and radiative energy and momentum output from stellar winds and supernovae (e.g. [Heckman et al. 1990]). They have been invoked as a source for the heating and metal-enrichment of both the intra-cluster and inter-galactic medium ([Adelberger et al. 2003; Loewenstein 2004], as a major driver of galaxy evolution (e.g. [Croton et al. 2006]), and as playing a crucial role in establishing the mass-metallicity relation among galaxies ([Tremonti et al. 2004]). Superwinds have also been hypothesized to be the mechanism through which a dust enshrouded ultraluminous in-
frared galaxy (ULIRG) transitions to an optically identified quasar (e.g. Sanders et al. 1988).

The dynamical evolution of a starburst-driven outflow has been extensively discussed (e.g. Chevalier et al. 1985; Suchkov et al. 1994; Wang 1995; Tenorio-Tagle & Munóz-Tuñon 1994; Strickland et al. 2000). Superwinds are generated when the kinetic energy in the outflow from massive stars and supernovae is thermalized, generating a region of very hot \((T \sim 10^8 \text{ K})\) low-density gas in the ISM of a starburst galaxy (Chevalier et al. 1985). This hot gas will expand, sweeping-up and entraining ambient gas and creating a hot cavity within the overall structure of the ISM. This region of hot gas is highly over-pressurized relative to its surrounding ambient ISM and will expand along the steepest pressure gradients. As it expands, it sweeps up and shocks ambient material, creating an outflowing “superwind” (Castor et al. 1973; Weaver et al. 1977). Contributing to this overall energy driven outflow is the radiation pressure on dust grains that are likely coupled to the ionized gas within the flow (Murray et al. 2003).

There is a significant amount of morphological, physical and kinematic evidence for the existence of superwinds in nearby starburst and infrared luminous galaxies (e.g. Heckman et al. 1990). Large-scale optical emission-line and associated X-ray nebulae are ubiquitous in starbursts (Heckman et al. 1990; Lehnert et al. 1999; Strickland et al. 2004), and these increase in size and luminosity from dwarf starbursts to ULIRGs, where they can be tens of kpc in size (Grimes et al. 2003). In addition, the winds are observed in molecular lines such as CO (Walter et al. 2002; Bolatto et al. 2013) and blue-shifted absorption features (e.g. NaD) seen projected against the starburst nuclei (Heckman et al. 2004; Rupke & Veilleux 2005; Martin 2003) indicating that the neutral, cold ISM is being swept-up in the outflow. Recently, vigorous outflows have also been detected in ULIRGs and LIRGs in molecular absorption lines with Herschel – PACS which suggest very high mass outflow rates (e.g. Sturm et al. 2011; Fischer et al. 2014; Veilleux et al. 2013; Spoon et al. 2013). ULIRGs which have higher AGN bolometric fractions appear to have higher terminal velocities and shorter gas depletion timescales than those with lower AGN bolometric fractions, suggesting that at least in ULIRGs, the AGN power is related to the feedback on the cold molecular gas (Veilleux et al. 2013; Cicone et al. 2014).

M82 is the brightest, nearest, and best-studied starburst galaxy with an outflow (Heckman et al. 1990; Shopbell et al. 1993; Lehnert et al. 1999). At a distance of about 3.5 Mpc (Jacobs et al. 2009), M82 affords an unparalleled opportunity to study a dusty outflow in great detail. The outflow in M82 has extensive and complex extra-planar polycyclic aromatic hydrocarbon (PAH) and H\(_2\) emission, evident from the Spitzer – IRAC and Spitzer – IRS observations presented in Engelbracht et al. (2006). Extensive PAH emission is also seen at 3.3\(\mu\)m in the AKARI data (Yamagishi et al. 2012). PAHs are thought to be the origin of strong, broad emission features arising in the mid-infrared spectra of star forming galaxies (e.g. Brandl et al. 2006; Smith et al. 2007a). M82 has also been observed in the [CII] and [O I] emission lines by Contursi et al. (2013) who find a low velocity outflow in the atomic gas. Complex extra-planar warm H\(_2\) knots and filaments extending more than \(\sim 5\) kpc above and below the galactic plane of M82 have been imaged in the near-infrared using narrow-band filters by Veilleux et al. (2009). Using wide-field imaging, they found that H\(_2\) is widespread, suppressed relative to PAHs in the fainter regions, and concluded that H\(_2\) is not a dynamically important component of the outflow. However, Veilleux et al. (2009) could only measure the very hot \((\sim 1000\text{ K})\) H\(_2\) gas, and the mass of very hot H\(_2\) gas is much smaller than the mass of the H\(_2\) gas probed by this study, as we show in Section 3.3.

In this paper, we study the M82 outflow with Spitzer – IRS. With Spitzer/IRS spectroscopy, we can study in detail all the PAH features, derive the grain properties through analysis of the PAH ratios, and compare these to the warm \((200\text{K})\) molecular gas and the ionized gas (via the strong Ne lines) as a function of position above and below the plane of M82 for the first time. The rotational lines of H\(_2\), observable in the infrared, unlike the near-infrared ro-vibrational lines, probe most of the warm molecular gas mass. We then use the ratio of pure rotational H\(_2\) lines to constrain the physical and excitation conditions of the gas. We use PAH 11.3/7.7 and 6.2/7.7 ratio maps to study PAH sizes and ionization across the outflow, [NeIII]/[NeII] to study the variation of the hardness of the ionization field, H\(_2\)/PAH to diagnose the origin of the H\(_2\) emission, and the H\(_2\)/[NeII] and PAH/[NeII] line flux ratios to help understand the survivability of the PAH grains in the wind. For the first time we can make robust maps of these ratios throughout the outflow of M82, and compare all of these to the spatially-resolved X-ray emission to understand the properties of the dust and molecular gas and ionized gas in the wind.

The paper is organized as follows: in § 2 we describe the observations and the data reduction; in § 3 we present and analyze the the spectral maps built from the observations, using the PAH 11.3/7.7 and 6.2/7.7, [NeIII]/[NeII], H\(_2\)/PAH, H\(_2\)/[NeII], and PAH/[NeII] line ratios, and also H\(_2\) excitation diagrams to diagnose the properties of the ionized gas, PAHs and the H\(_2\) emission; in § 4 we diagnose the origin of H\(_2\) emission using X-ray images and shock excitation models; and in § 5 we present the conclusions.

## 2 OBSERVATIONS AND DATA REDUCTION

The outflow of M82 was observed with Spitzer-IRS between November 10–14 2006 and May 9–22 2008, using all low-resolution modules. In Figure 1 we show the footprints of the IRS observations on the outflow of M82. We extracted only those with lower AGN bolometric fractions, suggesting that at least in ULIRGs, the AGN power is related to the feedback on the cold molecular gas (Veilleux et al. 2013; Cicone et al. 2014).

We used CUBISM (Smith et al. 2007a) to build SL and LL spectral cubes from the Spitzer-IRS Basic Calibration Data (BCDs). In order to create and subtract a local background from the wind spectra, we observed a region 20 kpc North of the M82 disk with both IRS modules. This region was observed with the same parameters (integration time per pixel and cycles per position) that were used for the wind. After verifying that no line, PAH or dust continuum was visible in the off field, we subtracted an average image
created from this region from the wind data at each location and for each IRS module before building the spectral cubes.

We used PAHFIT (Smith et al. 2007b) to create maps of spectral lines and PAH features. PAHFIT is a spectral fitting routine that decomposes IRS low-resolution spectra into broad PAH features, unresolved line emission, and continuum emission from dust grains. PAHFIT allows for the deblending of overlapping features, in particular the PAH emission, silicate absorption and fine-structure atomic, and warm H$_2$ emission lines. Because it performs a simultaneous fit to the emission features and the underlying continuum, PAHFIT works best if it can fit combined SL and LL spectra across the full 5 – 40µm range. In order to provide a stable long wavelength continuum from which to estimate the true PAH emission, including the broad feature wings, we formed a combined SL + LL cube. In order to investigate the spatial variations of the H$_2$/PAH ratio at the maximum resolution we decided to rotate and regrid the LL cubes to match the orientation and pixel size of the SL cubes (1″.85 × 1″.85). Thus we provide a stable long-term extrapolation of the LL fit to the SL fit. By calculating the flux in the new LL cubes using a bilinear interpolation (in surface brightness units), we assure that the flux over the native SL pixel scale is conserved.

Finally, we applied a small correction factor to each individual SL spectrum, necessary to match the continuum in the corresponding LL spectrum. This scale factor ranges between 0.65 – 0.99, and was calculated by comparing pixels in the spectral overlap region between SL and LL at approximately 14µm. This scale factor is applied before running PAHFIT on the combined SL + LL cubes. Such scaling is commonly observed in spatially extended sources as a result of the different entrance slit dimensions of the SL and LL spectrograph modules. In Figure 2 we present an example of the spectral decomposition performed at every location of the map with PAHFIT. The output of PAHFIT is saved for each feature, and is used to create all the maps (PAH, H$_2$, [NeII] and [NeIII]) at the native resolution of the SL data as shown in Figure 3. Although care was taken to avoid the FIR-bright nucleus of M82 during the spectral mapping, the LL1 data exhibit a “ripple” pattern that varies spatially and spectrally, and the LL2 data have a slightly elevated continuum level in some pixels (compared to SL1). We believe both of these to be the effect of light scattered into the nuclear IRS entrance slits. Therefore we have opted not to use the LL1 data in the fitting and analysis, except for the derivation of an upper limit on the H$_2$ S(0) line flux, calculated as the 3 – σ noise level at 28µm. We have excluded regions of the LL2 data cubes where the implied scale factor in the continuum between SL1 and LL2 is greater than 2.0.

3 RESULTS

3.1 Morphology of the dust, ionized gas and warm H$_2$ emission

In Figure 3 we present the total PAH emission map calculated by summing over all the PAH bands between 5 – 14µm, the [NeII]12.8µm flux map derived from the SL data cubes, the total H$_2$ map derived from the SL and LL2 cubes, and the [NeIII]15.6µm emission map. We find widespread emission with significant structure in all of the line maps. To the North, both PAH and [NeII] exhibit two prominent, broad radial filaments. To the South, the PAH map shows a complex structure with a bubble, or loop. This loop is also seen in the [NeII] map. The overall morphologies of the warm H$_2$ and H$_\alpha$ are generally very similar, as seen in the NIR by
Veilleux et al. (2009). Maps of the cold molecular gas made in CO (J=1-0) by Salak et al. (2013) are at lower spatial resolution, but also show extended, extraplanar gas.

We can study the physical properties of the ionized gas and PAHs by mapping the ratios between ionized lines such as [NeIII]/[NeII] and between the different PAH bands. If the ionized gas is photoionized, the Ne line ratio is sensitive to the hardness of the radiation field and the ionization parameter. The Ne ratios observed in the outflow are standard values for the hardness of the radiation field found in starburst galaxies. The ratio of the 6.2, 7.7 and 11.3 micron PAH features can be used to estimate the ionization state and the size of the small dust grains (e.g. Draine & Li 2001). Even the overall absence of PAH emission may be related to the ionization state of the gas since we do not expect gas emitting strongly in [NeII] to show much PAH emission, as the small grains are easily destroyed by energetic photons (e.g. Engelbracht et al. 2006; Madden et al. 2006; Wu et al. 2006). The 11.3µm/7.7µm ratio is high when PAHs are more neutral and larger, while the 6.2µm/7.7µm PAH ratio is high.
when PAH grains are smaller. However, we caution that variations in the PAH line ratios are also influenced by the structural properties and elemental abundances in small particles giving rise to the PAH features (e.g. Yamagishi et al. 2012).

In Figure 4, we present maps of the PAH 11.3/7.7 and 6.2/7.7 ratios, the H$_2$/PAH ratio and the [NeIII]/[NeII] ratio. The H$_2$/PAH ratio map was produced using a H$_2$ [S(1)-S(3)] total flux map and PAH (7.7 + 8.6µm) flux map. Although there is a great deal of structure in the PAH ratio maps, the general trend is that the 11.3/7.7 and the 6.2/7.7 ratios increase as we move away from the starburst disk. The PAH 11.3/7.7 is enhanced by a factor of $\sim$ 4 and the PAH 6.2/7.7 is enhanced by a factor of $\sim$ 3 at the northern edge of the outflow compared to the regions near the center of the starburst disk. There is an enhancement of the 11.3/7.7 and 6.2/7.7 ratios by $\sim$ 2 in the southeast edges of the outflow.

The H$_2$/PAH ratio varies between 0.01 – 1, and the [NeIII]/[NeII] ratio varies between 0.2 – 0.9.
between the PAH emission “fingers” (as particularly evident in the Northern outflow region in Figure 3), both maps show an enhancement of the ratio.

In Figure 4 we also mark the 12 regions from which we have extracted 5 – 20μm spectra and performed multi-temperature fits to the H$_2$ lines. We show these spectra in Figure 5. These regions were chosen to represent the full range of H$_2$/PAH ratios and to approximately sample the full spatial extent of the outflow. For comparison, we also show an integrated spectrum of the center of M82 (Beirão et al. 2008), which we use as a reference. In addition to the 12 regions, we also show a composite spectrum of the entire IRS map in Figure 5. We measure the fluxes of the PAH features between 5 – 14μm, the fine-structure lines [NeII] and [NeIII], and the rotational excitation lines of H$_2$ from S(0) to S(3) and present the results in Tables 1 and 2. Regions 1, 5, 6, 7 and 9 are the most similar to the reference Beirão et al. (2008). The spectrum labeled “Total” is the sum of all pixels in the map.

Figure 5. Mid-infrared 5 – 20μm spectra from selected regions in the outflow, as seen in Figure 3 and the center of M82 from Beirão et al. (2008). The spectrum labeled “Total” is the sum of all pixels in the map.

3.2 H$_2$/PAH variations in the M82 wind

In Figure 6 we present a plot of the PAH 11.3/7.7 ratio as a function of the 6.2/7.7 PAH ratio. The range of 6.2/7.7 PAH ratios in the outflow is much larger and the values are, on average, smaller than the disk. The outflow exhibits 6.2/7.7 and 11.3/7.7 ratios ranging between ~ 0.2 – 0.4. On average, the bulk of the superwind shows larger, more ionized grains than in the starburst disk, although there are regions where the ratios overlap. As an example of the different PAH ratios found in the outflow, we present in the upper right corner of Figure 6 two spectra from different regions of the outflow, and an average spectrum of the disk. In general, the outflow has stronger 17μm PAH and H$_2$ emission than the SB disk, and the spectrum in the region 2 kpc north of the disk (the blue spectrum shown in Figure 6) has higher 17μm feature and H$_2$ S(1) flux, indicating the prevalence of larger PAH grains and enhanced H$_2$ emission further away from the disk. In summary, we see a much larger range in the PAH ratios in the outflow than is seen in the starburst disk. The average spectrum of the wind has smaller 6.2/7.7 ratios and larger 11.3/7.7 ratios than the average spectrum of the starburst, suggesting the grains may be larger and more ionized, on average, than those in the starburst disk.

The range of H$_2$/PAH varies significantly as a function of distance from the starburst. This ratio has been previously used by Ogle et al. (2010) and Guillard et al. (2012) to show the relative importance of the UV versus mechanical heating of the gas. We can see that the H$_2$/PAH ratio decreases with increasing [NeII] flux. The vast majority of the mapped regions lie well above the values seen in star-forming galaxies in the local Universe (the dashed lines - see Roussel et al. (2007)). The H$_2$/PAH ratio values span two orders-of-magnitude, 0.01-1. Most of the M82 wind points have H$_2$/PAH between 0.01 and 0.10, but ~ 4% have extreme values, H$_2$/PAH > 0.5. The fact that the H$_2$/PAH ratios seen in the wind are well above those seen in the SB disk, the normal galaxies observed by Roussel et al. (2007), and PDR models of Guillard et al. (2012) (see section 4) suggests that the H$_2$ is not solely excited by stellar radiation from the starburst.

We can also investigate possible correlations of the H$_2$/PAH ratio with diagnostics of the atomic gas and the size and ionization state of the grains. In Figure 8 we show H$_2$/PAH as a function of the PAH 11.3/7.7 ratio, the PAH 6.2/7.7 ratio, the H$_2$ S(3)/S(1) ratio, and the [NeIII]/[NeII] ratio. We see that in the disk the H$_2$/PAH ratio increases with increasing PAH 11.3/7.7 ratio. While the trend of increasing H$_2$/PAH in the wind is similar to that seen in the disk, the wind values extend to much larger values of both
11.3/7.7 PAH and H$_2$/PAH than seen in the starburst. The scatter to high H$_2$/PAH and a much larger range in the 6.2/7.7 ratio is seen in the upper right panel of Figure 8 as well. In the lower left, we observe an anti-correlation between H$_2$/PAH and S(3)/S(1) in the disk as warmer gas has lower H$_2$/PAH. The wind points seem consistent with this trend but the uncertainties are large. This general lack of very warm molecular gas and high H$_2$/PAH ratios in the wind may be a function of inability of the MIR diagnostic features to probe the hottest gas. The [NeIII]/[NeII] also increases with H$_2$/PAH in the disk, and the wind points to the North seem consistent with the trend, but clearly there are many wind points in a few regions in the northern outflow which have much higher H$_2$/PAH ratios than seen in the disk at a given [NeIII]/[NeII] value.

To further help us understand what drives the variations of the H$_2$/PAH flux ratios, we present maps of H$_2$/[NeII] and PAH/[NeII] in Figure 9. In general, the regions with high H$_2$/PAH correspond to high PAH/[NeII], suggesting that it is not simply a decreasing PAH that causes the rise in H$_2$/PAH off the plane of M82. Only in the northern edge of the map do we see a lower PAH/[NeII] ratio, which might indicate reduced PAH heating efficiency. It is important to note, however, that while the PAH/[NeII] ratio does change by factors of a few across the map, this is much less than the variation observed in H$_2$/PAH. Therefore the latter cannot be solely explained by a change in the strength of the PAH emission alone.

Figure 6. Plot of the PAH ratio 11.3/7.7 as a function of 6.2/7.7 for 10$^9 \times 10^9$ regions in the outflow. Red represents the northern outflow, cyan the southern outflow and green represents the central 1 kpc in M82 from Beirão et al. (2008). The dashed lines represent neutral and ionized PAHs from Draine & Li (2001). The dot-dashed line represent the PAH ratio values for a PAH grain of 250 carbon atoms. The red and cyan crosses represent typical uncertainties in the north and southern outflows. The three spectra on the upper right corner are derived from different regions in the outflow (in units of MJy/sr). The blue spectrum is derived from a region 2 kpc north of the disk and is represented by the blue data points, the red spectrum is an average spectrum of the whole northern outflow, and the green spectrum the average spectrum of the central region.

Figure 7. Plot of the H$_2$ S(1)-S(3)/PAH (7.7+8.6μm) as a function of the [NeII] surface brightness. Each datapoint represents a 10$^9 \times 10^9$ region. Red points are the northern outflow, cyan the southern outflow, and green the central 1 kpc of M82 from Beirão et al. (2008). The dashed line and the dash-dotted lines represent the average and the 1σ variation, respectively, of the H$_2$ S(1)-S(3)/PAH (7.7+8.6μm) ratio in the star-forming galaxy sample of Roussel et al. (2007). The dotted-dashed line represents the upper limit in the H$_2$ S(1)-S(3)/PAH (7.7+8.6μm) ratio predicted by PDR models (Guillard et al. 2012). The error bar in red represent typical uncertainties in the northern outflow, while the typical uncertainties in the data for the southern outflow region are smaller than the points.

3.3 H$_2$ masses and temperatures

With the rotational H$_2$ emission line fluxes in Table 2, we can estimate the gas temperature, column density, and mass distributions, assuming that the ortho/para ratio is in local thermal equilibrium (LTE). In Figure 10 we plot the excitation diagrams for all regions, except 2 and 9, which did not have sufficient signal-to-noise for fitting the data with these sorts of models. Using the Boltzmann equation, we fit the H$_2$ lines with two components, each with different temperatures. We then derive the total column density, H$_2$ temperatures and mass distributions. In fitting the data we set the minimum allowed temperature to 100 K and treated every upper limit as a 3-sigma detection. The S(1) line is the line with the highest signal-to-noise. Therefore, when we move the S(0) line as a detection, we obtain essentially the same results.

The diagrams for Regions 1, 3, 4, 5, 7 and 11 look very similar in shape, derived from very similar temperatures for each component. In Region 8, both components have similar temperatures, while in Region 10, there is a large difference in temperature between the two components. We note that Regions 8 and 10 have the most discrepant S(0) upper limits relative to the other fits, making the mass estimates particularly uncertain. In Table 3 we present the properties of the warm H$_2$ gas in these regions of interest. In all the regions the temperature of the warmer component is above 500 K, especially warm when compared with typical star forming regions (e.g. Roussel et al. 2003). The column density of the bulk of the warm H$_2$ gas ranges by a factor of ~30. The mass ratio between the warmer and colder component varies between 0.003 – 0.02, with the lowest value observed in Region 8 and the highest in Region 1. The region with the most
luminous H$_2$ emission is Region 8, where the warmer component accounts for $\sim 95\%$ of the luminosity, whereas the next most luminous region, 11, the colder component accounts for $\sim 65\%$ of the luminosity.

We also integrated the emission over the entire observed nebula surrounding M82 and measured the line fluxes, which are included in Table 2. The total warm H$_2$ mass derived from the integrated emission is $\sim 4.8 \times 10^5$ M$_\odot$. For completeness, we note that Veilleux et al. (2004) finds a mass of H$_2$ at a temperature T$>$1000 K (thus hotter than our “hot” gas) of $1.2 \times 10^4$ M$_\odot$, based on near-infrared observations. The relatively low mass is not surprising since the near-infrared ro-vibrational lines are probing very high excitation gas which emits very efficiently in these lines. The mass of warm molecular gas in the wind that we are detecting with the IRS is thus more than a factor of 100 more than is estimated from the NIR line maps. We have also estimated the H$_2$ masses from fitting the H$_2$ line fluxes with molecular shock models (see Section 4.3).

Considering a cold molecular gas mass of $\sim 10^9$M$_\odot$ (Walter et al. 2002; Salak et al. 2013), we calculate a ratio of warm to cold molecular gas in the outflow of about 0.01. Therefore the warm molecular gas we measure here is a relatively minor component of the total molecular mass in the ISM. Care must be taken when making this comparison as the emission regions for the CO and H$_2$ do not overlap completely. In fact, the most extended CO emission is only within a few arc minutes of the plane and much less extended (Salak et al. 2013; Walter et al. 2002), although new studies trace the CO-emitting gas to larger distances (Leroy et al. 2014). However, we note that most of the estimated mass for the warm molecular gas in our observations also comes from regions which lie close to the disk (and likely the base of the wind).

4 DISCUSSION

4.1 H$_2$ heating processes

In the scenario of thermal excitation of H$_2$, three mechanisms dominate the heating: (1) UV radiation from starburst, (2) X-rays from the starburst and wind, and/or (3) shocks induced by the outflow. Following Ogle et al. (2010), we use the H$_2$-to-PAH luminosity ratio to investigate the contribution of UV photons to the total heating of the H$_2$.
gas in the outflow. In starburst galaxies like M82, we expect the H$_2$ emission to be predominantly powered by UV radiation from young stars but undoubtedly, processes such as low-velocity shocks and turbulent dissipation contribute to the H$_2$ emission. It should also be noted that H$_2$/PAH ratios on the order of 0.01 have also been seen from translucent clouds in the cold, diffuse ISM of the Milky Way, which are thought to be excited by UV photons (Ingalls et al. 2011). X-rays due to inverse Compton scattering of electrons (Begelman et al. 1989), high velocity shocks in the expanding bubble of hot and tenuous gas generated by the intense star formation within M82 (Ogle et al. 2010). Therefore, we conclude that X-rays are not the dominant source of heating of the H$_2$ gas in the M82 wind, since the observed ratios are well above those found in XDR models.

There is a clear offset of the wind points from those in the SB disk in terms of H$_2$, PAH and ionized atomic emission (Figs. 3, 8, and 9). The data in regions of the outflowing gas show a large range in both [NeIII]/[NeII] and H$_2$/PAH, but are clearly offset to higher H$_2$/PAH at a given [NeIII]/[NeII] ratio compared to regions within the SB disk, suggesting an additional excitation mechanism at work. In fact, models suggest that the large values of the H$_2$/PAH ratios observed in some regions (>0.1) cannot be explained by photo-ionization (Guillard et al. 2012).

The H$_2$ gas could be indirectly excited by X-rays emitted from the expanding bubble of hot and tenuous gas generated by the intense star formation within M82 (Begelman et al. 1989). High velocity shocks in the entrained gas, X-rays due to inverse Compton scattering of IR photons. In M82, the X-ray emission covers approximately the same region as the optical emission line filaments (Watson et al. 1984; Shopbell et al. 1998) and radio emission (Scaquist & Odegard 1991). For example, the high H$_2$/PAH in the northern tip ~5 kpc from the disk, appears to be associated with a region with X-ray emission (Figure 11).

We can use the total H$_2$S(1) − S(3) luminosity and the total X-ray luminosity in the outflow to calculate the H$_2$ to-X-ray luminosity ratio (Stevens et al. 2003). The H$_2$ transition can occur at relatively low column densities, N_H ≈10$^{21}$ cm$^{-2}$. At such low columns, even relatively soft X-rays can penetrate. Our estimates for the M82 outflow are L(H$_2$S(0) − S(3))/L_X (2–10 keV) ~ 1 (Stevens et al. 2003) and L(H$_2$S(0) − S(3))/L_X (0.2–4 keV) ~ 0.5 (Watson et al. 1984; Strickland & Heckman 2007), over the northern and southern IRS map. Using XDR models (e.g. Maloney et al. 1996), we can estimate whether the H$_2$ line emission could be powered by X-ray heating, following the procedure described in Ogle et al. (2010). In these models, 30% − 40% of the absorbed X-ray flux goes into gas heating via photoelectrons, and the cooling by H$_2$ rotational lines in XDR models is ~ 2% of the total gas cooling for a gas temperature of 200 K. At this temperature, the ratio of the first four rotational lines to the total rotational line luminosity is L(H$_2$ 0–0 S(0)−S(3))/L_X = 0.6. Combining the above factors, we estimate a maximum H$_2$ to-X-ray luminosity ratio of L(H$_2$ 0–0 S(0)−S(3))/L_X (2–10 keV) < 0.01 (Guillard et al. 2012). This upper limit on the H$_2$-to-X-ray luminosity is conservative because it assumes that all the X-ray flux is absorbed by the molecular gas, which is likely not the case (since the medium in which X-rays propagate is inhomogeneous). Therefore, we conclude that X-rays are not the dominant source of heating of the H$_2$ gas in the M82 wind, since the observed ratios are well above those found in XDR models. This is also seen in some extreme environments, such as powerful radio galaxies, where H$_2$/PAH ~ 0.5 − 1 and the H$_2$/X-ray ratio also exceeds that expected from XDRs alone (Ogle et al. 2010).

Ambient gas clouds heated by the mechanical energy of the wind and dense shell fragments from the wind bubble carried along by the wind are the two main sources of X-ray emission in the wind (Strickland et al. 1997). In Figure 11, we estimate a maximum H$_2$ to-X-ray luminosity ratio of L(H$_2$ 0–0 S(0)−S(3))/L_X (2–10 keV) < 0.01 (Guillard et al. 2012). This upper limit on the H$_2$-to-X-ray luminosity is conservative because it assumes that all the X-ray flux is absorbed by the molecular gas, which is likely not the case (since the medium in which X-rays propagate is inhomogeneous). Therefore, we conclude that X-rays are not the dominant source of heating of the H$_2$ gas in the M82 wind, since the observed ratios are well above those found in XDR models. This is also seen in some extreme environments, such as powerful radio galaxies, where H$_2$/PAH ~ 0.5 − 1 and the H$_2$/X-ray ratio also exceeds that expected from XDRs alone (Ogle et al. 2010).

Ambient gas clouds heated by the mechanical energy of the wind and dense shell fragments from the wind bubble carried along by the wind are the two main sources of X-ray emission in the wind (Strickland et al. 1997). In Figure 11, we estimate a maximum H$_2$ to-X-ray luminosity ratio of L(H$_2$ 0–0 S(0)−S(3))/L_X (2–10 keV) < 0.01 (Guillard et al. 2012). This upper limit on the H$_2$-to-X-ray luminosity is conservative because it assumes that all the X-ray flux is absorbed by the molecular gas, which is likely not the case (since the medium in which X-rays propagate is inhomogeneous). Therefore, we conclude that X-rays are not the dominant source of heating of the H$_2$ gas in the M82 wind, since the observed ratios are well above those found in XDR models. This is also seen in some extreme environments, such as powerful radio galaxies, where H$_2$/PAH ~ 0.5 − 1 and the H$_2$/X-ray ratio also exceeds that expected from XDRs alone (Ogle et al. 2010).

Ambient gas clouds heated by the mechanical energy of the wind and dense shell fragments from the wind bubble carried along by the wind are the two main sources of X-ray emission in the wind (Strickland et al. 1997). In Figure 11, we estimate a maximum H$_2$ to-X-ray luminosity ratio of L(H$_2$ 0–0 S(0)−S(3))/L_X (2–10 keV) < 0.01 (Guillard et al. 2012). This upper limit on the H$_2$-to-X-ray luminosity is conservative because it assumes that all the X-ray flux is absorbed by the molecular gas, which is likely not the case (since the medium in which X-rays propagate is inhomogeneous). Therefore, we conclude that X-rays are not the dominant source of heating of the H$_2$ gas in the M82 wind, since the observed ratios are well above those found in XDR models. This is also seen in some extreme environments, such as powerful radio galaxies, where H$_2$/PAH ~ 0.5 − 1 and the H$_2$/X-ray ratio also exceeds that expected from XDRs alone (Ogle et al. 2010).

Figure 9. Maps of the H$_2$S(1) − S(3)/[NeII] ratio (left), PAH5 − 17µm/[NeII] (right). H$_2$ is the sum of the S(1), S(2), and S(3) H$_2$ line fluxes and PAH5 − 17µm is the sum of the fluxes of all PAH features between ~ 5 and 17µm. All maps have been clipped at the 3 − σ surface brightness level. The white radial strip in the bottom two panels are regions that have been masked out due to scattered light in the LL2 module, as in Figures 11.
Figure 10. $H_2$ excitation diagrams of most selected regions. The dotted lines show a 2-temperature LTE fit of the data. The upper limits of the $H_2$ lines represent a 3-σ noise level at the corresponding line.
the increase of H$_2$/PAH(7.7+8.6) surface brightness. The morphology of H$_2$/PAH closely resembles the morphology of the H$_2$ (2.12µm)/PAH maps in [Veilleux et al. 2008]. The regions where H$_2$/PAH > 0.1 correspond to regions where H$_2$ (2.12µm)/PAH > 0.003. While there is a reasonable spatial correlation between the surface brightness distribution of the soft X-ray and H$_2$ emission, a correlation is apparent between soft X-rays surface brightness and H$_2$/PAH ratio distributions only in high surface brightness X-ray emitting regions. There are several possibilities why this might be so, but the most obvious is that both the X-ray and H$_2$ emission are fundamentally powered by the same mechanism, namely, the very hot plasma that is acting as the “piston” driving the wind and heating clouds entrained in the wind. However, the warm H$_2$ and hot X-ray emitting plasma should not be co-spatial on small scales, and instead the H$_2$ is likely seen in projection against the hot wind, along the edges of the cone in dense knots of gas, heated by shocks (see below). Such a situation would also arise in gas with a wide range of densities or through changes as individual clouds are shredded as they are accelerated by the action of the piston [Cooper et al. 2008]. This is also shown by the strong correlation between the X-rays and the PAH emission (Figure 11), especially in the Southern wind. As the PAHs and H$_2$ molecules cannot be expected to survive within the $10^6-10^7$ K X-ray emitting plasma, it is most likely that they lie along the edges of this phase of the outflow. So while the power source is likely the same, the two emission regions, not surprisingly, represent different physical regions and gas states being impacted by the driving very hot plasma [Strickland & Heckman 2007].

Figure 11. Image of the H$_2$/PAH (top) and total PAH flux (5.2–17µm) (bottom), overlaid with contours from soft (0.3–1.6 keV) X-ray emission map of M82 (Strickland & Heckman 2009).

we show the contours of soft X-ray emission from the outflow of M82 overlaid on the H$_2$/PAH and total (5–17µm) PAH maps. The X-ray and PAH maps show a remarkable correlation, particularly in the southern part of the bi-conical outflow, although both maps are projections of the true 3D structure and the precise locations of the hot gas and dust grains are unknown. Moreover, there are regions of high surface brightness X-ray emission with relatively little PAH emission (in the north along the outer wind axis). We also see that the middle lobe of X-ray emission in the northern outflow, at about 5 kpc from the disk, appears related to

4.2 The velocities and kinetic energy of the molecular gas

Kinematics of the gas can be used to determine if the different gas phases observed in the outflowing wind are related. If the observed kinematics of the warm molecular gas as probed by the H$_2$ rotational lines is similar to that of the cold molecular gas as probed by the CO lines, the IR atomic neutral and ionized emission lines, or the lines of the warm ionized medium, then this might suggest a direct physical relationship between these phases. [Salak et al. (2013)] derived a velocity of 160 km/s for the CO(1-0) emitting gas for the northern outflow region. However, the farthest extent of the CO emission is only about 1 arc minute north of the disk of M82, while the IRS map extends to 4 arcmin. The velocity derived by [Salak et al. 2013] is not very different from the velocities in the neutral and ionized atomic gas in the infrared within the inner wind region < 1 arcmin from the disk [Contursi et al. 2013]. If we restrict ourselves to the inner wind region within a few arc minutes of the plane of M82, the velocities in the warm ionized gas are only a few 100 km s$^{-1}$ or less within a couple of arc minutes of the plane of M82 [Shopbell et al. 1998]. These velocities are consistent with those observed in IR neutral and ionized atomic lines [Contursi et al. 2013]. However, farther out of the plane of M82, much higher outflow velocities (∼600–800 km s$^{-1}$) are observed in H$_2$ for the outflow in M82 [Shopbell et al. 1998]. Line splitting is observed in both the CO and H$_2$ with similar velocities [Walter et al. 2002], making the association between the cold molecular
Figure 12. Shock excitation diagrams of most selected regions in the outflow. The best fit models for either J- (light blue) or C-shocks (dark blue) or a combination of each are indicated in the legend to each panel (see text for details). The sum of or the individual the best fit models is shown as a solid red line. The level populations of the various rotational levels of the (0-0) H$_2$ lines are indicated by plus signs if they are detection or downward pointing arrows for upper limits and are labeled as S(1), S(2), S(3), S(4), and S(5).
gas and the warm ionized gas stronger (cf. Salak et al. 2013; Shopbell et al. 1998).

Such a direct relationship would be consistent with a picture where the emission we observe in the extended outflow of M82 is mainly from entrained and swept-up clouds that are predominately shock-heated by a high ram-pressure, hot plasma generated within the starburst region. If we envision the regions of warm H2 emission as being the areas in the clouds that are shock-heated by the expanding plasma, we would expect these phases to co-exist with the warm ionized gas (Heckman et al. 1990) and the cold/warm plasma, we would expect these phases to co-exist with the areas in the clouds that are shock-heated by the expanding plasma, we would expect these phases to co-exist with the warm ionized gas (Heckman et al. 1990) and the cold/warm molecular gas.

Assuming that the velocity of the warm H2 is similar to the deprojected velocity of the cold molecular gas derived by Salak et al. (2013), we can calculate its total kinetic energy as $E_K = 0.5 \times M_{H_2} v_{\text{outflow}}^2 \approx 5.9 \times 10^{54}$ erg for $M_{H_2} \sim 4.8 \times 10^6 M_\odot$ and $v_{\text{outflow}} \approx 160$ km s$^{-1}$. Our estimates are considerably higher than those of Veilleux et al. (2002) because both their mass estimates are a factor of 500 less and they assumed a velocity of 100 km s$^{-1}$ for the warm molecular emission. However, they are smaller compared to those estimated for the cold molecular and warm ionized medium. Salak et al. (2013) estimate that the kinetic energy of the wind in the cold molecular gas is $\sim 1 \times 4 \times 10^{56}$ erg. Walter et al. (2002) estimate a kinetic energy of $\approx 3 \times 10^{55}$ erg, but they do not map the CO distribution out to large distances from the disk compared to Salak et al. (2013).

### 4.3 H$_2$ supersonic turbulent heating and shock models

The hot plasma driving the outflow will likely transfer part of its ram and thermal pressure to swept-up and entrained dense molecular gas by driving shocks and by turbulent mixing between hot and cold gas (see Guillard et al. 2007). To model such shocks, we use a grid of stationary shock models computed with an updated version of the Lesaffre et al. (2013) code to derive the physical shock parameters that explain the H$_2$ line ratios. Our initial conditions are that the gas is diffuse, magnetized with a pre-shock magnetic field of intensity $B/\mu G = \sqrt{n_H/\text{cm}^{-3}}$, and irradiated with a radiation field of $G_0 = 1$ in Habing units. As the “Habing field” would photodissociate the H$_2$, we assume some self-shielding. In the models, we integrate the extinction along the model from the pre-shock, and we take into account the shielding and self-shielding, which is parametrized in the photo-dissociation reaction rates of H$_2$ and CO (Lesaffre et al. 2013). At a density of $10^4$ cm$^{-3}$, the ionization fraction in the pre-shock gas is $6 \times 10^{-8}$. This matches approximately our estimate of the ionization fraction of $10^{-7}$ at 10 kpc from the X-ray source. We fit the H$_2$ line fluxes observed in the different regions with one shock or a combination of two shock velocities depending on how many H$_2$ lines are detected, for three different pre-shock densities: $n_H = 10^3$, $10^4$ and $10^5$ cm$^{-3}$, (see Guillard et al. 2004 for details of the method). The results of these fits are shown in Figure 12 where we present the H$_2$ excitation diagrams for the best velocity combinations (lowest $\chi^2$), and for the intermediate pre-shock density of $n_H = 10^3$ cm$^{-3}$. We present in Table 4 the results of the shock models for each region and the integrated spectrum (labeled “Total”). The total gas mass derived from the shock models is $M_{H_2} = 1.7 \times 10^7 M_\odot$, which results in an estimate of the kinetic energy of $E_K \sim 2 \times 10^{54}$ erg. This is a factor of 3 higher than the estimates calculated in the temperature diagrams of Figure 13. This difference arises due to the fact that we are integrating the postshock temperature profile down to 50 K (where the H$_2$ mid-IR lines start to emit), and because local thermodynamic equilibrium (LTE) is not assumed in the shock models.

For all regions, the best-fit shock velocities are in the range 4-35 km s$^{-1}$ (grids spans 3-40 km s$^{-1}$; Table 4). These velocities are much lower than the estimated bulk outflow velocity. Lehnert et al. (1999). The regions show a range of best fitting models, but typically the best fitting pre-shock densities are $n_H = 100-1000$ cm$^{-3}$. Our modeling is not intended to be exhaustive, but to simply provide an estimate of the physical parameters in the warm molecular gas of the outflow.

At the range of preshock densities and magnetic field strengths considered, the C-shocks turn into J-shocks when the velocities are above $22$ km s$^{-1}$. For some regions the H$_2$ line fluxes are reproduced by a combination of two C-shocks, in other regions the C-shocks (J-shocks) do not sufficiently populate the highest (lowest) energy levels. In those cases, we fit the H$_2$ line fluxes by a combination of C- and J-shocks. The shocks travel through the clouds in $\sim 10^4 - 10^5$ years (for 1 pc clouds), which is much shorter than the dynamical time of the outflow ($10^5$ years - Lehnert et al. (1999)).

Using the estimates of the warm molecular gas mass and the typical cooling time of the shocked H$_2$ gas (t$_{cool} \sim 10^4$ yr for $n_H = 10^3$ cm$^{-3}$), we can estimate the mass flow of shocked gas, $M_{\text{warm}}/t_{\text{cool}} \sim 7.5 \times 10^6 M_\odot/10^4$ yr $= 750 M_\odot$ yr$^{-1}$, assuming that all the warm H$_2$ is shock-heated. The dynamical timescale of the H$_2$-rich outflow is $\sim 4.5$ kpc/200 km s$^{-1} = 2.5 \times 10^7$ yr implying a total mass of shocked gas $\approx 5 \times 10^3 M_\odot$. This estimate is likely to be an upper limit since the H$_2$ gas may not extend over the whole outflow and because this assumes that the gas is only shocked once, which is highly unlikely. The high mass flow rate of shocked gas may require the reformation of H$_2$ in the post-shocked gas (see Guillard et al. 2002, 2010) to prolong the survival time of the H$_2$ gas in the outflow (e.g. Fragile et al. 2004).

Clouds may also survive longer if the radiative cooling time is shorter than the time scale necessary for the shocks to cross the cloud and the growth rate of Kelvin-Helmholtz instabilities (Cooper et al. 2008). Such processes would lower the total molecular mass required to explain the molecular emission from the outflow, although for the C shocks considered here and even for many of the J shocks as well, there is little destruction of H$_2$. At any rate, the total gas mass that is shock heated is similar to that of the cold molecular gas. Therefore the reservoir for the shocked gas is likely the cold molecular gas as probed through the CO emission from the wind.

### 4.4 PAH destruction

Since we see that the PAHs grains are on average larger and more ionized in the wind, it is natural to ask whether or not some of the enhanced H$_2$/PAH could be do to a reduction in the PAH emission, rather than an enhancement of H$_2$. The decrease in PAH flux is probably just due to a decrease of the radiation field, but there is also a change
in the grain properties associated with the wind. The gradient of decrease in H$_2$ emission is significantly lower than that of the PAH emission, suggesting an additional heating source for the H$_2$ (see Section 3.2). One of the ways PAH emission can decrease relative to H$_2$ is PAH destruction. PAH grains can be destroyed either by hard radiation field, X-rays, or shocks. The EWs of the PAH features in starburst galaxies are observed to decrease sharply when $[\text{NeIII}]15\mu\text{m}/[\text{NeII}]12.8\mu\text{m}>2$, where $[\text{NeIII}]/[\text{NeII}]$ traces the hardness of the radiation field. This has been interpreted as evidence of PAH destruction by intense radiation fields (Wu et al. 2006; Beirão et al. 2006). In M82, despite an increase of $[\text{NeIII}]15.6\mu\text{m}/[\text{NeII}]12.8\mu\text{m}$ in the outflow, in most of the outflow $[\text{NeIII}]15.6\mu\text{m}/[\text{NeII}]12.8\mu\text{m}$ remains less than 1. $[\text{NeIII}]15.6\mu\text{m}/[\text{NeII}]12.8\mu\text{m}$~1 indicates that there is hard (>40.96 eV) radiation present, but not enough to have an effect on the average PAH properties. Furthermore, only a small fraction of the gas in the M82 outflow is photoionized, and PAHs may survive in HI and H$_2$ gas. PAH destruction is also very sensitive to the clumping factor of the gas. If the gas is dense and clumpy with relatively small volume filling factor and large covering fraction, the PAH grains will be shielded from the hard photons in the radiation field. So PAH destruction by hard photons is not likely to be the cause of the increase of H$_2$/PAH in the wind.

PAH grains are rapidly destroyed in hot, X-ray emitting gas (Micelotta et al. 2010b). However, the presence of PAHs at large distances (10 kpc from the disk) in the outflow shows that they are protected in cooler phases of the gas, mostly in molecular gas. It is possible, however, that X-rays from the hot outflow fluid could destroy the PAH carriers, but the correlation between the X-ray emitting gas and the PAH maps, particularly in the Southern outflow, suggests that grain destruction by X-rays in the wind is not dominant.

Thermal collisions with very small and big grains can have a significant impact on the cooling of the hot gas ejected into the halo (Dwek et al. 1990). However, the impact of this cooling will generally be short lived since the cooling efficiency drops as the grain sputtering occurs in the hot gas (Guillard et al. 2009), unless the hot halo gas is constantly replenished with dusty gas. A physical dust cooling model (Guillard et al. 2009), unless the hot halo gas is constantly replenished with dusty gas. A physical dust cooling model (Guillard et al. 2009), unless the hot halo gas is constantly replenished with dusty gas. A physical dust cooling model (Guillard et al. 2009), unless the hot halo gas is constantly replenished with dusty gas. A physical dust cooling model (Guillard et al. 2009), unless the hot halo gas is constantly replenished with dusty gas. A physical dust cooling model (Guillard et al. 2009), unless the hot halo gas is constantly replenished with dusty gas. A physical dust cooling model (Guillard et al. 2009), unless the hot halo gas is constantly replenished with dusty gas. A physical dust cooling model (Guillard et al. 2009), unless the hot halo gas is constantly replenished with dusty gas.

Another possible mechanism for PAH destruction is shocks. As seen in Figure 12, the H$_2$ emission that we see can be fit with shock models having speeds of up to 40 km s$^{-1}$. PAH destruction likely only becomes relevant for shock velocities >100 km s$^{-1}$ (Micelotta et al. 2010a). It is unlikely that the low velocity shocks that heat the molecular gas and enhance the H$_2$ emission are also destroying the PAH grains, although we clearly see a change in the grain population. Similarly, Contursi et al. (2013) find that shocks do not dominate the [CII] and [OI] emission in the wind. Overall, the enhancement in H$_2$/PAH is likely caused by H$_2$ excitation by shocks, although we cannot rule out a relatively small effect due to PAH destruction.

In dense molecular clouds, where the PAH destruction by shocks is less efficient, cosmic rays are the most efficient source of destruction of PAH molecules, mostly because of their bombardment by energetic ions (Micelotta et al. 2011).

This is likely to play an important role in the M82 wind, because the intense star formation in M82 is likely to enhance the cosmic ray intensity by a factor of $\approx 5$ compared to the Milky Way (Robitaille et al. 2010). Theoretical studies suggest that CRs can destroy PAHs in dense molecular clouds, but in these environments the heating will also be very inefficient, causing a reduction of PAH emission. Micelotta et al. (2011) showed that only PAHs with a number of carbon atoms greater than $\approx 250$ have a survival lifetime greater than the circulation time scale of the outflow ($\approx 2 \times 10^7$ yr). Thus, there may be a size-selection effect of PAHs in the outflow, which is roughly consistent with the size inferred from the band ratios (we mostly see PAH 6.2/7.7 < 0.3, meaning an average size greater than 250 carbon atoms, see Figure 6), and cosmic rays are likely to play a role.

The calculations of Micelotta et al. 2010 show that any PAH shocked at $>100$ km/s should be destroyed. To be re-formed in the post-shock gas, an efficient and fast reformation process is needed. PAHs are a product of high temperature chemistry involving abundant carbon bearing precursors such as CH$_4$ and C$_2$H$_2$. In the cooled postshock gas these precursor species are likely not very abundant. In the cold neutral medium, the global PAH coagulation timescale is of the order of 100 Myr, much longer than the outflow dynamical time (Seok et al. 2014). AKARI results from Yamagishi et al. (2012) suggest that very small grains still exist in the halo. Therefore it is reasonable to assume that PAH are produced in the disk, carried up in the outflow, and "protected" from fast shocks in the denser material. If molecular clouds are formed in the postshock gas in the outflow, then there may be PAH coagulation over timescales comparable to the outflow time (coagulation timescales are of the order of a few Myr in molecular clouds), but it should be very small fraction of the volume, contrary to what is observed.

In conclusion, the enhancement of the H$_2$/PAH is likely caused by H$_2$ shock excitation, although we cannot rule out a relatively small effect due to PAH destruction.

5 CONCLUSION

- We have mapped the Northern and Southern superwind in the nearby starburst galaxy M82 using the Spitzer-IRS. We find strong H$_2$, PAH and Ne line emission that extends over nearly 5 arcminutes above and below the plane of the galaxy, and have measured the variation of the H$_2$/PAH, $[\text{NeIII}]/[\text{NeII}]$ and PAH feature ratios throughout the wind.

- We estimate a total warm (100 – 500 K) molecular gas mass to be between $\sim 5 - 17 \times 10^9$ M$_\odot$ and the total kinetic energy of the warm molecular outflow to be between $6 - 20 \times 10^{50}$ erg assuming an outflow velocity $v_{\text{outflow}}=160$ km s$^{-1}$. The warm molecular gas provides about $0.5-5\%$ of the kinetic energy – comparable to that observed in the atomic gas.

- There are clear variations in the PAH band ratios throughout the wind. The 6.2/7.7 and 11.3/7.7 micron PAH band flux ratios, when compared to models, suggest a higher fraction of larger and/or more ionized PAHs in the wind compared to the starburst disk, especially in the Northern part of the wind.

- The warm molecular gas to PAH ratio (H$_2$/PAH) is
enhanced in the outflow by factors of 10-100 as compared to the starburst disk. This enhancement in the $H_2/PAH$ ratio does not seem to follow the trends with the ionization of the atomic gas (as measured with the $[NeIII]/[NeII]$ line flux ratio) seen in the starburst. Our results indicate that UV and X-ray emission are probably not the dominant excitation mechanisms of $H_2$. Models suggest the observed values could be reproduced with slow C- and J-shocks ($v < 40$ km s$^{-1}$) driven into dense molecular gas entrained in the outflow. Using a dynamical timescale of the $H_2$-rich outflow of about $10^7$ years results in an upper limit of $5 \times 10^8 M_\odot$ of molecular gas heated and shocked over the lifetime of the wind. Given the rate at which the gas must be shocked, the reservoir for this shocked gas is the cold molecular gas as observed in CO.

* The emission from all of these phases is likely due to clouds entrained in and accelerated by the flow. They are being shocked heated by the hot plasma with high ram pressure generated by the collective thermalization of stellar winds and supernova in the starburst region of M82. These clouds are multiphase and the variation in properties are likely due to a range of shock velocities which depend on the density of the pre-shocked gas which depend on the density of the pre-shocked gas.

**REFERENCES**

Adebahr B., Krause M., Klein U., Wezgowicz M., Romans D. J., & Dettmar R.-J. 2013, A&A, 555, 23
Adelberger, K. L., Steidel, C. C., Shapley, A. E., & Pettini, M. 2003, ApJ, 584, 45
Appleton, P. N., Xu, K. C., Reach, W., et al. 2006, ApJ, 639, 51
Begelman, M. C., & Cioffi, D. F. 1989, ApJ, 345, 21
Beirão, P., Brandl, B. R., Appleton, P. N., Groves, B., Armus, L., Förster Schreiber, N. M., Smith, J. D., Charmandaris, V., & Houck, J. R. 2008, ApJ, 676, 304
Beirão, P., Brandl, B. R., Devost, D., Smith, J. D., Hao, L., & Houck, J. R. 2006, ApJ, 643, 1
Bolatto, A. D., Warren, S. R., Leroy, A. K. et al. 2013, Nature, 499, 450
Brandl, B. R., Bernard-Salas, J., Spoon, H. W. W., et al. 2006, ApJ, 653
Castor, J., McCray, R., & Weaver, R. 1975, ApJ, 200, 107
Chevalier, R. A. & Clegg, A. W. 1985, Natur., 317, 44
Chièze, J.-P., Pineau des Forêts, G., & Flower, D. R. 1998, MNRAS, 295, 672
Cicone, C., Maiolino, R., Sturm, E., et al. 2014, A&A, 562, 21
Contursi, A., Poglitsch, A., Garcia Carpio, J., et al. 2013, A&A, 549, 118
Cooper, J. L., Bicknell, G. V., Sutherland, R. S., & Blundell-Hawthorn, J. 2008, ApJ, 674, 157
Croton, D. J., et al. 2006, MNRAS, 365, 11
Draine, B. T., & Li, A. 2001, ApJ, 551, 807
Dwek, E., Rephaeli, Y., & Mather, J. C. 1990, ApJ, 350, 104
Elmegreen, B. G 1979, ApJ, 232, 729
Engelbracht, C. W., Kudrunth, P., Gordon, K. D. et al. 2006, ApJ, 642, 127
Fischer, J., Sturm, E., Gonzalez-Alfonso, E. et al. 2010, A&A, 518, 41
Fragile, P. C., Murray, S. D., Anninos, P., & van Breugel, W. 2004, ApJ, 604, 74
Grimes, J. P., Heckman, T., Strickland, D., & Ptak, A. 2005, ApJ, 628, 187
Guillard, P., Ogle, P. M., Emonts, B. H. C., Appleton, P. N., Morganti, R., Tadhunter, C., Oosterloo, T., Evans, D. A., & Evans, A. S. 2012, ApJ, 747, 95
Guillard, P., Boulanger, F., Pineau Des Forêts, G. et al. 2012, ApJ, 749, 158
Guillard, P., Boulanger, F., Chuver, M. E., Appleton, P. N., Pineau Des Forêts, G., & Ogle, P. 2010, A&A, 518, 59
Guillard, P., Boulanger, F., Pineau Des Forêts, G., & Appleton, P. N. 2009, A&A, 502, 515
Heckman, T. M., Armus, L., & Miley, G. K. 1990, ApJ, 74, 833
Heckman, T. M., Lehnerd, M. D., Strickland, D. K., & Armus, L. 2000, ApJS, 129, 493
Higdon, S. J. U., Armus, L., Higdon, J. L., Soifer, B. T., & Spoon, H. W. W. 2006, ApJ, 648, 323
Houck, J. R. 2008, ApJ, 676, 304
Ingalls, J. G., Bania, T. M., Boulanger, F., Draine, B. T., Falgarone, E., & Hily-Blant, P. 2011, ApJ, 743, 174
Jacobs, B. A., Rizzi, L., Tully, R. B., Shaya, E. J., Makarov, D. I. & Makarova, L. 1999, AJ, 138, 322
Klein, R. I., McKee, C. F., & Colella, P. 1994, ApJ, 420, 213
Leeuw, L. L., & Robson, E. I. 2009, AJ, 137, 517
Lehnerd, M. D., Heckman, T. M., & Weaver, K. A. 1999, ApJ, 523, 575
Leroy, A. et al. 2014, in prep.
Lesaffre, P., Pineau des Forêts, G., Godard, B., Guillard, P., Boulanger, F., & Falgarone, E. 2013, 550, 106
Lesaffre, P., Chièze, J.-P., Cabrit, S., & Pineau Des Forêts, G. 2004, A&A, 427, 147
Loewenstein, M. 2004, in Origin and Evolution of the Elements, Carnegie Observatories Centennial Symposia, ed. A. McWilliam & M. Rauch (Cambridge University Press), 2004, 422
Madden, S. C.; Galliano, F.; Jones, A. P., & Sauvage, M. 2006, A&A, 446, 877
Maloney, P. R., Hollenbach, D. J., & Tielens, A. G. M. 1996, ApJ, 466, 561
Martin, C. L. 2005, ApJ, 621, 227
Micelotta, E. R., Jones, A. P., & Tielens, A. G. M. 2011, A&A, 526, 52
Micelotta, E. R., Jones, A. P., & Tielens, A. G. M. 2010, A&A, 510, A36
Micelotta, E. R., Jones, A. P., & Tielens, A. G. M. 2010, A&A, 510, A37
Murray, N., Quataert, E., & Thompson, T. A. 2005, ApJ, 618, 569
Ogle, P., Boulanger, F., Guillard, P., Evans, D. A., Antonucci, R., Appleton, P. N., Nesvadba, N., & Leipski, C. 2010, ApJ, 724, 1193
Rigopoulou, D., Kunze, D., Lutz, D., Genzel, R., & Moorwood, A. F. M. 2002, A&A, 389, 374
Robitaille, T. P., & Whitney, B. A. 2010, ApJ, 710, L11
Rousset, H., Helou, G., Hollenbach, D. J., et al. 2007, ApJ, 669, 959
Rupke, D. S., & Veilleux, S. 2005, ApJ, L37
Salak, D., Nakai, N., Miyamoto, Y., Yamauchi, A., & Tsuru, T. G. 2013, PASJ, 65, 66
Sanders, D. B., Soifer, B. T., Elias, J. H., Neugebauer, G., & Matthews, K. 1988, ApJ, 328, 35
Sanders, D. B., Soifer, B. T., Elias, J. H., Neugebauer, G., Matthews, K. 1988, ApJ, 328, 35
Seaquist, E. R., & Odegard, N. 1991, ApJ, 369, 320
Shopbell, P. L., & Bland-Hawthorn, J. 1998, ApJ, 493, 129
Smith, J. D. T., Armus, L., Dale, D. A., Roussel, H., Sheth, K., Buckalew, B. A., Jarrett, T. H., Helou, G., & Kennicutt, R. C., Jr. 2007, PASP, 119, 1133
Smith, J. D. T., Draine, B. T., Dale, D. A. et al. 2007, ApJ, 656, 770
Spoon, H. W. W., Farrah, D., Lebouteiller, V., et al. 2013, ApJ, 775, 127
Stevens, I. R., Read, A. M., & Bravo-Guerrero, J. 2003, MNRAS, 343, 47
Strickland, D. K., & Heckman, T. M. 2007, ApJ, 658, 258
Strickland, D. K., Heckman, T. M., Colbert, E. J. M., Hoopes, C. G., & Weaver, K. A. 2004, ApJ, 606, 829
Strickland D.K., Ponman T.J., & Stevens I.R. 1997, A&A, 320, 378
Strickland, D. K., & Stevens, I. R. 2000, MNRAS, 314, 511
Strickland, D. K., & Stevens, I. R. 2000, MNRAS, 314, 511
Sturmi, E. et al. 2011, ApJ, 733, 16
Suchkov, A. A., Balsara, D. S., Heckman, T. M., & Leitherer, C. 1994, ApJ, 430, 511
Seok, J. Y., Hirashita, H. & Asano, R. S. 2014, MNRAS, 439, 2186
Tenorio-Tagle, G., Silich, S., & Muñoz-Tuñón, C. 2003, ApJ, 597, 279
Tenorio-Tagle, G., & Muñoz-Tuñón, C. 1998, MNRAS, 293, 299
Tremonti, C. A., Heckman, T. M., Kauffmann, G. et al. 2004, ApJ, 613, 898
de Vaucouleurs, G., de Vaucouleurs, A., Corwin, H. G., Jr., Buta, R. J., Paturel, G., & Fouqu, P. 1991, Third Reference Catalogue of Bright Galaxies, Vols. 13 (Berlin: Springer)
Veilleux, S., Melndez, M., Sturm, E. et al. 2013, ApJ, 776, 27
Veilleux, S., Rupke, D. S. N., & Swaters, R. 2009, ApJ, 700, 149
Walter, F., Weiss, A., & Scoville, N. 2002, ApJ, 580, 21
Walmsley, M., Pineau des Forêts, G., & Flower, D. 2005, in Astrochemistry: Recent Successes and Current Challenges, eds. D. C. Lis, G. A. Blake, & E. Herbst, IAU Symp., 231, 135
Wang, B. 1995, ApJ, 444, 590
Watson, M. G., Stanger, V., & Griffiths, R. E. 1984, ApJ, 286, 144
Weaver, R., McCray, R., Castor, J., Shapiro, P., & Moore, R. 1977, ApJ, 218, 377
Wu, Y., Charmandaris, V., Hao, L., Brandl, B. R., Bernard-Salas, J., Spoon, H. W. W., & Houck, J. R. 2006, ApJ, 639, 157
Yamagishi, M., Kaneda, H., Ishihara, D., Kondo, T., Onaka, T., Suzuki, T., & Minh, Y. C. 2012, A&A, 541, 10
Table 1. PAH band and ionic line fluxes (in units of $10^{-16}$ Wm$^{-2}$) and ratios. The regions for the extracted spectra are shown in Figure 3 and described in the text.

| Region | 6.2µm | 7.7µm | 8.6µm | 11.3µm | 17µm | PAH 6.2 – 12.6µm | 6.2/7.7 | 11.3/7.7 | [NeII] | [NeIII] |
|--------|-------|-------|-------|--------|------|------------------|--------|--------|-------|--------|
| 1      | 7.12  | 23.31 | 4.19  | 6.44   | 3.52 | 44.58            | 0.305  | 0.151  | 3.62  | 1.36   |
| 2      | 1.86  | 3.31  | 1.28  | 2.56   | 0.920| 9.93             | 0.562  | 0.278  | 2.60  | 0.966  |
| 3      | 1.01  | 2.58  | 0.526 | 1.12   | 0.751| 5.987            | 0.391  | 0.291  | 1.27  | 0.550  |
| 4      | 0.811 | 4.76  | 0.895 | 2.65   | 1.76 | 10.876           | 0.170  | 0.370  | 0.536 | 0.893  |
| 5      | 12.53 | 40.41 | 8.93  | 13.86  | 8.32 | 84.05            | 0.310  | 0.206  | 2.17  | 5.38   |
| 6      | 7.30  | 22.03 | 6.04  | 10.65  | 7.94 | 53.96            | 0.331  | 0.360  | 2.79  | 4.18   |
| 7      | 9.39  | 29.26 | 6.61  | 11.43  | 7.47 | 64.16            | 0.321  | 0.255  | 6.06  | 3.44   |
| 8      | 349.6 | 1167.7| 254.3 | 397.7  | 282.6| 2451.9           | 0.299  | 0.242  | 339.3 | 89.4   |
| 9      | 741.9 | 2783.7| 513.5 | 668.9  | 254.3| 4962.3           | 0.261  | 0.091  | 568.0 | 83.1   |
| 10     | 10.3  | 37.6  | 82.7  | 18.3   | 10.7 | 159.6            | 0.274  | 0.285  | 5.43  | 2.57   |
| 11     | 30.9  | 105.7 | 22.2  | 37.0   | 21.3 | 217.1            | 0.292  | 0.202  | 30.2  | 9.91   |
| 12     | 8.22  | 27.2  | 7.55  | 15.0   | 9.40 | 67.37            | 0.302  | 0.346  | 9.37  | 4.17   |
Table 2. $\text{H}_2$ line fluxes (in units of $10^{-17}$ Wm$^{-2}$) and ratios. "Total" denotes the sum of all emission in the maps.

| Region | $\text{H}_2$ S(0) | $\text{H}_2$ S(1) | $\text{H}_2$ S(2) | $\text{H}_2$ S(3) | $\text{H}_2$ S(1) - S(3) |
|--------|------------------|------------------|------------------|------------------|------------------|
| 1      | $<0.313$        | $4.13\pm0.12$   | $0.840\pm0.142$ | $1.70\pm0.14$   | $6.67$           |
| 2      | $<0.169$        | $2.81\pm0.07$   | $0.517\pm0.267$ | $1.68\pm0.23$   | $5.01$           |
| 3      | $<1.45$         | $6.68\pm0.14$   | $1.74\pm0.16$   | $3.17\pm0.14$   | $11.6$           |
| 4      | $<0.991$        | $14.3\pm0.2$    | $3.93\pm0.24$   | $8.74\pm0.41$   | $27.0$           |
| 5      | $<2.81$         | $9.68\pm0.26$   | $6.24\pm0.55$   | $2.05\pm0.18$   | $18.0$           |
| 6      | $<1.25$         | $8.87\pm0.14$   | $2.37\pm0.14$   | $4.79\pm0.20$   | $16.0$           |
| 7      | $<48.7$         | $37.8\pm0.6$    | $32.2\pm0.4$    | $83.4\pm0.7$    | $153.4$          |
| 8      | $<101.9$        | $14.3\pm0.2$    | $84.6\pm0.3$    | $108.4\pm0.3$   | $372.8$          |
| 9      | $<528$          | $23.9\pm0.2$    | $5.73\pm0.11$   | $119.0\pm3.0$   | $148.6$          |
| 10     | $<51$           | $2.883$         | $2.12$          | $4.28\pm1.47$   | $1.43$           |
| 11     | $<531$          | $2.945$         | $0.838$         | $12.9\pm10.1$   | $7.91$           |
| 12     | $<657$          | $2.940$         | $2.07$          | $4.19\pm0.82$   | $2.26$           |
| Total  | $<253.9$        | $1171.5\pm3.2$  | $455.1\pm2.7$   | $1010.0\pm5.6$  | $2636.6$         |

Table 3. $\text{H}_2$ Temperature Model Parameters. "Total" denotes the sum of all selected regions.

| Regions | T(K)   | Ortho/Para x10$^{19}$ cm$^{-2}$ | $N(H_2)$ x10$^4$M$_\odot$ | $M(H_2)$ x10$^4$L$_\odot$ | $L(H_2)$ |
|---------|--------|---------------------------------|----------------------------|----------------------------|----------|
| Region1 | 210. ± 18. | 2.883                           | 2.12                       | 4.28±1.47                  | 1.43     |
|         | 1094. ± 208. | 3.000                           | 0.05                       | 0.0107±0.0036              | 1.83     |
| Region3 | 240±54   | 2.945                           | 0.838                      | 12.9±10.1                  | 7.91     |
|         | 819±126  | 3.000                           | 0.007                      | 0.113±0.087                | 7.90     |
| Region4 | 236. ± 14. | 2.940                           | 2.07                       | 4.19±0.82                  | 2.26     |
|         | 831. ± 124. | 3.000                           | 0.0128                     | 0.03±0.01                  | 1.84     |
| Region5 | 206. ± 17. | 2.8871                          | 6.77                       | 13.7±4.1                   | 4.32     |
|         | 549. ± 54. | 3.000                           | 0.132                      | 0.27±0.11                  | 4.37     |
| Region6 | 166. ± 21. | 2.653                           | 13.5                       | 27.3±18.2                  | 3.25     |
|         | 593. ± 56. | 3.000                           | 0.086                      | 0.152±0.052                | 3.30     |
| Region7 | 224. ± 12. | 2.917                           | 3.30                       | 6.66±1.17                  | 2.96     |
|         | 657. ± 75. | 3.000                           | 0.0382                     | 0.0771±0.0297              | 2.68     |
| Region8 | 324.±12.  | 2.994                           | 5.32                       | 10.7±0.6                   | 64.0     |
|         | 595.±13.  | 3.000                           | 0.80                       | 1.6±0.2                    | 3.6      |
| Region10| 218.±6.   | 2.913                           | 9.5                        | 19.2±1.7                   | 33.0     |
|         | 637.±62.  | 3.000                           | 0.091                      | 0.18±0.06                  | 5.2      |
| Region11| 233. ± 3.  | 2.934                           | 11.4                       | 23.0±0.81                  | 11.8     |
|         | 623. ± 14. | 3.000                           | 0.02                       | 0.39±0.03                  | 10.1     |
| Region12| 241. ± 10. | 2.946                           | 2.61                       | 5.28±0.666                | 3.17     |
|         | 816. ± 157. | 3.000                           | 0.02                       | 0.042±0.002                | 3.09     |
| Total   | 272±2     | 2.976                           | 1.39                       | 476±7.66                  | 463      |
|         | 854±18    | 3.000                           | 0.023                      | 7.67±0.42                 | 640      |
Table 4. H$_2$ Shock Model Parameters. "Total" denotes the sum of all selected regions.

| Regions | $v_{\text{shock}}$ km s$^{-1}$ | $M_{\text{flow}}$ $M_\odot$/yr | $t_{\text{cool}}$ $\times 10^3$ yr | $M(H_2)$ $\times 10^4 M_\odot$ | Type |
|---------|-------------------------------|---------------------------------|---------------------------------|---------------------------------|------|
| Region1 | 5.5                           | 42±3.1                          | 1.6                             | 6.8±0.5                         | C    |
|         | 35.0                          | 16±1.4                          | 3.4                             | 5.4±0.47                        | J    |
| Region3 | 5.5                           | 210±35                          | 1.6                             | 34±5.6                          | C    |
|         | 33.0                          | 150±11                          | 3.2                             | 48±3.5                          | J    |
| Region4 | 6.0                           | 48±6.1                          | 1.6                             | 7.6±0.96                        | C    |
|         | 35.0                          | 16±2.4                          | 3.4                             | 5.5±0.83                        | J    |
| Region5 | 5.0                           | 200±12                          | 1.6                             | 33±2.0                          | C    |
|         | 20.0                          | 1.9±0.14                        | 2.1                             | 0.41±0.03                       | C    |
| Region6 | 4.0                           | 1200±200                        | 1.7                             | 200±33                          | C    |
|         | 21.5                          | 1.3±0.31                        | 2.1                             | 0.27±0.06                       | C    |
| Region7 | 5.0                           | 130±32                          | 1.6                             | 21±5.2                          | C    |
|         | 21.5                          | 0.97±0.14                       | 2.1                             | 0.2±0.029                       | C    |
| Region8 | 5.5                           | 400±91                          | 1.6                             | 64±15                           | C    |
|         | 21.5                          | 17±4.4                          | 2.1                             | 3.6±0.93                        | C    |
| Region10| 5.5                           | 210±40                          | 1.6                             | 33±6.3                          | C    |
|         | 21.5                          | 1.5±0.24                        | 2.1                             | 0.31±0.05                       | C    |
| Region11| 5.0                           | 500±63                          | 1.6                             | 80±10                           | C    |
|         | 18.0                          | 6.2±0.26                        | 2.1                             | 1.3±0.05                        | C    |
| Region12| 5.5                           | 84±6.3                          | 1.6                             | 13±0.98                         | C    |
|         | 22.0                          | 33±1.4                          | 2.2                             | 7.5±0.32                        | J    |
| Total   | 5.5                           | 10000±2100                      | 1.6                             | 1700±360                        | C    |
|         | 21.5                          | 180±14                          | 2.1                             | 39±3.1                          | C    |