CONSTRAINTS FROM \(^{26}\text{Al}\) MEASUREMENTS ON THE GALAXY'S RECENT GLOBAL STAR FORMATION RATE AND CORE-COLLAPSE SUPERNOVAE RATE

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ABSTRACT

Gamma rays from the decay of \(^{26}\text{Al}\) offer a stringent constraint on the Galaxy's global star formation rate over the past million years, supplementing other methods for quantifying the recent Galactic star formation rate, such as equivalent widths of Hz emission. Advantages and disadvantages of using \(^{26}\text{Al}\) gamma-ray measurements as a tracer of the massive star formation rate are analyzed. Estimates of the Galactic \(^{26}\text{Al}\) mass derived from COMPTEL measurements are coupled with a simple, analytical model of the \(^{26}\text{Al}\) injection rate from massive stars and restrict the Galaxy's recent star formation rate to \(5 \pm 4 \, \text{M}_\odot\, \text{yr}^{-1}\). In addition, we show that the derived \(^{26}\text{Al}\) mass implies a present-day Type II + Ib supernovae rate of \(3.4 \pm 2.8\) per century, which seems consistent with other independent estimates of the Galactic core-collapse supernova rate. If some independent measure of the massive star initial mass function or star formation rate or Type II + Ib supernovae rate were to become available (perhaps through estimates of the Galactic \(^{60}\text{Fe}\) mass), then a convenient way to restrain, or possibly determine, the other parameters is presented.

Subject headings: Galaxy: stellar content — gamma rays: theory — nuclear reactions, nucleosynthesis, abundances — stars: formation — stars: statistics — supernovae: general

1. INTRODUCTION

Through its 1.809 MeV gamma-ray line, Galactic \(^{26}\text{Al}\) was discovered in 1979 with the High Energy Astronomy Observatory C spectrometer (Mahoney et al. 1982). Several measurements of the integrated 1.809 MeV flux have been performed since then, as reviewed by Prantzos & Diehl (1996). The most reliable of these measurements are probably derived from data obtained from the Gamma-Ray Spectrometer aboard the Solar Maximum Mission spacecraft (Share et al. 1985; Harris, Share, & Leising 1994), and the COMPTEL Imaging Telescope aboard the Compton Gamma-Ray Observatory (Diehl et al. 1995a). All estimates of the absolute \(^{26}\text{Al}\) mass in the Galaxy rest on assumptions about the spatial distribution of the sources, as the 1.809 MeV measurements themselves do not carry distance information. This situation may change with future high-resolution, high-sensitivity instruments if the line shape and Doppler shift of the 1.809 MeV line can be extracted (Gehrels & Chen 1996).

From the HEAO C data, Mahoney et al. determined \(3 \, \text{M}_\odot\) of \(^{26}\text{Al}\), assuming the smooth spatial distribution derived from COS B measurements of Galactic gamma rays in the 100 MeV regime. All other nonimaging measurements and analysis can be interpreted as 0.7 of \(^{26}\text{Al}\) if only the spiral-arm component is assigned to Type II + Ib supernovae, to 2.5 \(\text{M}_\odot\) of \(^{26}\text{Al}\) if the emission is assigned to Type II/Ib supernovae and no foreground contributions from localized emission regions similar to Vela or Cygnus lie in the direction of the inner Galaxy. Since the decay time of \(^{26}\text{Al}\) \((\tau_{1/2} = 7.5 \times 10^5 \, \text{yr})\) is short compared to Galactic rotation timescales \((\tau_{\text{gal}} \approx 10^8 \, \text{yr})\), this estimated \(^{26}\text{Al}\) mass range serves as an important
constraint of the stellar population responsible (i.e., massive stars) for the synthesis.

One of the recurring concepts of this paper is the relationship between the derived $^{26}$Al mass estimates, initial mass function (IMF), recent global star formation rate (SFR), and present epoch Type II + 1b supernovae rate. The connections made in this paper between these four quantities are new, or at least not widely recognized, but the raw science of the four quantities themselves is not new. It is worthwhile in making the connections to succinctly summarize the shape of the IMF, tracers of the Galactic SFR, and stellar core-collapse estimates.

2. STAR FORMATION RATE CONSTRAINTS

2.1. The Initial Mass Function

The first empirical determination of the observed IMF showed that the number of stars between 0.4 and 10 $M_\odot$ could be described as a power law with the index $\gamma = -2.35$ (Salpeter 1955). Studies since then (Miller & Scalo 1979; Humphreys & McElroy 1984; Scalo 1986; Rana 1991; Tinney, Mould, & Reid 1992; Parker & Garmany 1993; Reid 1994; Hunter 1995; Massey et al. 1995a, 1995b; Kroupa 1995; Méra, Chabrier, & Baraffe 1996; Mayya & Prabhu 1996; Hunter et al. 1996) suggest that the observed IMF becomes flatter than a pure power law at the smallest stellar masses ($\gamma \approx -1$ for $m \leq 0.5 M_\odot$) and becomes steeper for the most massive stars ($\gamma \approx -3.3$ for $m > 10 M_\odot$). Some studies indicate that the IMF has more structure than either a power law or log-normal form (Rana 1991), while others argue that the IMF is closer to a power law and has less structure (Scalo 1986). Overall, the shape of the IMF appears to be quite robust (centered on the Salpeter $-2.35$ exponent) and seems not to change very much from one star-forming region to another. In any event, proceeding from the observed luminosity function to the implied IMF depends upon the stellar evolutionary tracks used in the fitting procedure (Tinsley 1980; Elmegreen 1995a, 1995b; Efremov 1995; Adams & Fatuzzo 1996; Arnett 1996).

2.2. Galactic Star Formation Rates

Despite many uncertainties, the results of various studies (Schmidt 1959, 1963; Searle, Sargent, & Bagnuolo 1973; Larson & Tinsley 1974; Cohen 1977; Smith, Biermann, & Mezger 1978; Lequeux 1979; Talbot 1980; Tinsley 1980; Kennicutt 1983; Kennicutt & Kent 1983; Güsten & Mezger 1983; Turner 1984; Lacey & Fall 1985; Dopita 1985; Gallagher, Bushouse, & Hunter 1989; Romanishin 1990; Rana 1991; Kennicutt 1992; Lada 1992; Gallagher & Gibson 1993; Kennicutt, Tantyl, & Condgon 1994; Hill, Madore, & Freedman 1994; Gallagher & Scowen 1995; Gallego et al. 1995; Lada & Lada 1995) lead to the picture of $\sim 10\%$ of the current SFR occurring in the innermost 1 kpc of the Galaxy, and most of the remaining 90% concentrated between 5 and 9 kpc from the center, which is where most of the Galaxy’s giant molecular clouds, infrared emission, and other signs of intense star formation reside. The current SFR for the whole Galaxy has been estimated to be $0.8 M_\odot$ yr$^{-1}$ (Talbot 1980), $3.0 M_\odot$ yr$^{-1}$ (Turner 1984), $5.3 M_\odot$ yr$^{-1}$ (Smith et al. 1978), $13.0 M_\odot$ yr$^{-1}$ (Güsten & Mezger 1982), and $6.0 M_\odot$ yr$^{-1}$ (Page 1994). Güsten & Mezger (1982) also estimated the massive star SFR in the spiral arms to be $5 \pm 2 M_\odot$ yr$^{-1}$. Note that all of these estimates use various indicators of the massive star population, and then convert these indicators to a total SFR by means of an assumed (universal) IMF.

2.3. Hα Line Widths

Equivalent widths of Hα emission have been the most available and most popular method for quantifying the present SFR. This is because Hα equivalent widths are directly proportional to the number of Lyman continuum photons emitted by massive stars and hence proportional to the SFR. Other measures or indicators (HI, Hα, [O III] λ5007, [O II] λ3727, integrated UBV colors, infrared luminosities, IRAS fluxes, free-free radio emission from H α regions, magnetic field strengths, and brightest individual star counts) are more affected by stellar absorption, interstellar reddening, excitation strength, metallicity, dust abundances, dust composition, incompleteness, sky coverage, or resolution limitations than Hα emission (see references above). Even when Hα measurements are combined with some of the alternative indicators, the derived SFR is a lower limit since even Hα is not completely immune (just less sensitive) from the contaminants listed above. It has been suggested that near-infrared recombination lines of Brγ could be an even better measure of the current SFR (Leitherer & Heckmann 1995), but instrumentation difficulties impede progress along this avenue at present.

2.4. Gamma-Ray Measurements

Gamma rays from the decay of $^{26}$Al offer a unique measure of the present SFR in the Galaxy. Several of the difficulties noted above are mitigated by the transparency of the Galaxy to gamma rays (e.g., absence of interstellar reddening), but several difficulties remain (e.g., spatial resolution limitations). Nevertheless, gamma rays offer a complementary indicator of the Galaxy’s present epoch SFR.

The IMF by number (assumed universal and constant independent), the normalization condition, and the normalization constant are

$$f(m) = \int_{M_L}^{M_U} f(m) dm = 1,$$

$$A = (\gamma + 1)(M^\gamma_{U} - M^\gamma_{L})^{-1}, \quad \gamma \neq -1,$$

respectively. The mean mass of stars, the fraction of all stars which become core-collapse supernovae (Type II + Type Ib), and the steady state core-collapse supernovae rate, respectively, are

$$\langle m \rangle = \int_{M_L}^{M_U} f(m)m dm = \frac{A}{\gamma + 2}(M_{U}^{\gamma+1} - M_{L}^{\gamma+1}) M_\odot,$$

$$F_{SN} = \int_{M_{SN}}^{M_U} f(m) dm = \left[ 1 - \left( \frac{M_{SN}}{M_U} \right)^{\gamma+1} \right] \left[ 1 - \left( \frac{M_{L}}{M_U} \right)^{\gamma+1} \right]^{-1},$$

$$R_{SN} = N_\star F_{SN} = \Psi \frac{F_{SN}}{\langle m \rangle} \text{ number yr}^{-1},$$

where $N_\star$ is the stellar birthrate in number per year, $\Psi$ is the SFR in $M_\odot$ per year, $M_L$ is the largest stellar mass in the distribution, $M_U$ is the smallest stellar mass in the distribution, and $M_{SN}$ is the smallest stellar mass which undergoes...
core collapse. The mean yield of $^{26}$Al from Type II + Ib events is

$$\langle y \rangle = \int_{M_{\text{SN}}}^{M_{\text{U}}} f(m) y(m) dm \left[ \int_{M_{\text{SN}}}^{M_{\text{U}}} f(m) dm \right]^{-1}$$

$$= \frac{1}{F_{\text{SN}}} \int_{M_{\text{SN}}}^{M_{\text{U}}} f(m) y(m) dm$$

$$= \gamma_{\text{eff}} M_{\odot} \cdot F_{\text{SN}}.$$  \hspace{1cm} (3)

Finally, the steady state injection rate of $^{26}$Al is

$$M_{26} = R_{\text{SN}} \langle y \rangle = \Psi \gamma_{\text{eff}} \langle m \rangle M_{\odot} \text{ yr}^{-1}.$$  \hspace{1cm} (4)

Equation (4) may be solved for the global SFR, $\Psi$, for a given observed $^{26}$Al mass in a steady state galaxy, and a given IMF exponent $\gamma$. The results of such a procedure is shown in the lower panel of Figure 1. Each labeled curve corresponds to a different Galactic $^{26}$Al mass (in solar masses), with the preferred $^{26}$Al mass range ($0.7$–$2.8$ $M_{\odot}$) imposed by the COMPTEL observations shown as the gray band. Integration limits of $M_{\odot} = 0.1 M_{\odot}$, $M_{\text{U}} = 40 M_{\odot}$, and $M_{\text{SN}} = 10 M_{\odot}$ were used in constructing Figure 1, but the chief conclusions are quite robust with respect to reasonable variations in the integration limits. The mean and effective $^{26}$Al yields in equations (3) and (4) were calculated with the $^{26}$Al mass ejected in the Woosley & Weaver (1995) massive star models. There is about 1 order of magnitude difference in the $^{26}$Al yields if the results of the Thielemann, Nomoto, & Hashimoto (1996) survey are used instead of Woosley & Weaver. The bulk of the synthesis of this radioactive isotope takes place in the presupernova star. It is imperative to follow this stage of the star’s evolution with a sufficient nuclear reaction network, especially during the last few hours of convective neon and oxygen burning. Woosley & Weaver used a 200 isotope network from the main sequence through the explosion, while Thielemann et al. follow the presupernova evolution from an initial helium core mass with an $\alpha$-chain network. Only during the explosive phases of the evolution do Thielemann et al. switch to a larger reaction network. This accounts for most of the difference in the $^{26}$Al production in the two surveys. Deviation from straight lines in the lower panel of Figure 1 is due to the IMF exponent approaching the removable singularity at $\gamma = -1$ (see eq. [1]). Only a mathematical reason, not a physical one, is responsible for the flattening of the curves.

The horizontal dimension of the dashed box in the lower panel of Figure 1 is centered on the Salpeter $–2.35$ exponent and is representative of the range of IMF exponents for massive stars encountered in the literature. Vertical dimensions of the dashed box were set by requiring consistency between the COMPTEL estimates of the Galactic $^{26}$Al mass and the simple model for the $^{26}$Al injection rate from massive stars. The lower panel of Figure 1 suggests that the global SFR in the Galaxy during the past million years is restricted to $5 \pm 4 M_{\odot}$ yr$^{-1}$. This is consistent with the Güsten & Mezger (1982) H$\alpha$ estimate of the massive star SFR in the spiral arms of $5 \pm 2 M_{\odot}$ yr$^{-1}$, and the more recent determinations of the Galaxy’s global SFR (see § 1).

Equation (2) may be solved for the core-collapse supernova rate given the global SFR $\Psi$ and he IMF exponent $\gamma$. This solution is shown in the upper panel of Figure 1, for the SFRs calculated in constructing the lower panel of Figure 1. As before each labeled curve corresponds to a different Galactic $^{26}$Al mass, with the preferred $^{26}$Al mass range ($0.7$–$2.8$ $M_{\odot}$) imposed by the COMPTEL observations shown as the gray band. Here the curves are straight lines (as expected) since the approach to the removable singularity at $\gamma = -1$ is embedded in both factors ($\langle m \rangle$, $F_{\text{SN}}$) of equation (2) and they cancel each other. For the same plausible range of IMF exponents considered above, the COMPTEL estimates of the Galactic $^{26}$Al mass appear to imply a core-collapse supernovae rate of $3.4 \pm 2.8$ per century.

**3. DISCUSSION**

Direct measurement of the Galactic supernova rate is difficult owing to possible incompleteness in historical observations and uncertainty as to the fraction of the Galactic disk and altitude that are sampled. Indirect inference from supernova rates in similar galaxies is adversely affected by the imprecise value of the Hubble constant and the uncertainty in estimating the total blue luminosity and morphological classification of our Galaxy. Systematic searches for extragalactic supernova are also hampered by
the need to know the distance, luminosity, and Hubble class of the host galaxy, as well as the dates and limiting magnitude of each observation. Such detailed information is available only in a few dozen supernova catalogs. Based on these surveys, estimates of the core-collapse and thermonuclear-driven supernova rates were derived and discussed by van den Bergh & Tammann (1991) and Cappellaro (1993). These estimates assumed that the peak luminosity of each supernova class was a standard candle, and a large correction for edge-on spirals (sin e effect). Using the extragalactic estimates with a total Galactic blue luminosity of 2.3 × 10^{10} L_{\odot}, a Hubble constant of 75 km s^{-1}Mpc^{-1}, and a Sbc Galactic morphology, the Galactic core-collapse supernova rate has been estimated to be 4.1 per century (van den Bergh & Tammann 1991) and 2.4–2.7 per century (van den Bergh & McClure 1994; Tammann, Löffler, & Schröder 1994). These estimates agree (perhaps auspiciously) with the core-collapse supernova rate implied by a near-Salpeter IMF exponent and the COMPTEL-derived 26Al mass.

While the general agreement found between estimates of the COMPTEL-derived 26Al mass, the range of massive star IMF exponents encountered in the literature, complementary measures of the recent SFR, and the present epoch Type II + Ib supernovae rate may be fortuitous and reminiscent of epicycles, it does point to a consistent picture. Less speculative is the fact that gamma rays from the decay of certain radioactive nuclei, such as 26Al and 60Fe, offer a unique measure of the present SFR in the Galaxy that is complimentary to other popular indicators of the Galaxy’s present epoch SFR (e.g., Hz, Hβ, Hγ, [O iii] λ5007, [O ii] λ3727, integrated UV colors, infrared and radio luminosities, and stellar counts). If some independent measure of the massive star IMF exponent or SFR or Type II + Ib supernovae rate were to become available (perhaps through measurements of the Galactic 60Fe mass), then Figure 1 offers a convenient way to constrain, or possibly even determine, the other parameters.

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