ABSTRACT

We analyze the BOOMERANG 2003 (B03) 145 GHz temperature map to constrain the amplitude of a non-Gaussian, primordial contribution to CMB fluctuations. We perform a pixel-space analysis restricted to a portion of the map chosen in view of high-sensitivity, very low foreground contamination and tight control of systematic effects. We set up an estimator based on the three Minkowski functionals which relies on high-quality simulated data, including non-Gaussian CMB maps. We find good agreement with the Gaussian hypothesis and derive the first limits based on BOOMERANG data for the nonlinear coupling parameter $f_{NL}$ as $-300 < f_{NL} < 650$ at 68% CL and $-800 < f_{NL} < 1050$ at 95% CL.

Subject heading: cosmic microwave background

1. INTRODUCTION

While cosmology is entering its precision era, the target of experiments aimed at the cosmic microwave background (CMB) is shifting toward weak signals, e.g., polarization, the SZ effect, and non-Gaussian (NG) features. Detection of NG signals can be of significant help in constraining the mechanisms that explain the generation of cosmological perturbations. Provided that systematic effects will not degrade high-sensitivity CMB mapping, present and future experiments could in principle be sensitive to nonlinearities due to second-order effects in perturbation theory (Bartolo et al. 2004). This signal is usually parameterized by a nonlinear coupling factor $f_{NL}$ that controls the level of a quadratic contribution to the primordial gravitational potential $\Phi$ (Komatsu & Spergel 2001):

$$\Phi(x) = \Phi_0(x) + f_{NL} [\Phi_0(x)^2 - \langle \Phi_0(x)^2 \rangle],$$

where $\Phi_0$ is a zero-mean, Gaussian random field.

Several groups have reported NG constraints on CMB data. All suborbital efforts to date have found no significant deviation from Gaussianity in the CMB field: MAXIMA-1 reported $|f_{NL}| < 950$ at 1 $\sigma$ (Santos et al. 2003; Wu et al. 2001), while the Very Small Array found an upper limit of 5400 at 2 $\sigma$ (Smith et al. 2004); the Archeops group have recently improved their limits to $-800 < f_{NL} < 1100$ (2 $\sigma$), although their analysis is based on assumptions only valid for the large angular scales dominated by the Sachs-Wolfe effect (Curto et al. 2007). The BOOMERANG 1998 data set has also been tested for Gaussianity, both in pixel (Polenta et al. 2002) and in Fourier (De Troia et al. 2003) space, finding no trace of NG signals. However, BOOMERANG has set no limit so far. One of the purposes of this Letter is to provide such limits with the analysis of the new 2003 (B03) data. The limits presented here are more stringent than those found by any suborbital experiment to date, properly accounting for subhorizon angular scales.

The WMAP team constrained $f_{NL}$ to be $-54 < f_{NL} < 114$ (Spergel et al. 2006). Using an improved version of the WMAP team estimator, Creminelli et al. (2007) set the most stringent limits to date at $-36 < f_{NL} < 100$. Thus the $f_{NL}$ analysis does not show any departure from Gaussianity in WMAP data. However, some authors have looked at general deviations from Gaussianity (i.e., not based on any specific parameterization of NG) and claimed highly significant detection of NG features in the WMAP maps (Copi et al. 2004; Vielva et al. 2004; Cruz et al. 2006).

In this Letter we perform a pixel-space analysis of the B03 temperature ($T$) data set, using the observed field’s moments and Minkowski functionals (MFs) to build Gaussianity tests. We assess the statistical significance of our results by com-
paring the data to a set of highly realistic, Gaussian Monte Carlo (MC) simulated maps. In order to constrain $f_{\text{NL}}$, we build a goodness-of-fit statistics based on MFs and calibrated against a set of NG CMB maps that are generated according to the algorithm set forth in Liguori et al. (2003).

The plan of this Letter is as follows: in § 2 we briefly describe the B03 experiment, the data set it has produced, and our simulation pipeline. In § 3 we compute the map’s moments and MFs of the data and compare results against Gaussian MC simulated maps. Then we derive constraints for $f_{\text{NL}}$. Finally, in § 4 we draw our main conclusions.

2. THE BOOMERANG 2003 DATA SET

The balloon-borne B03 experiment was flown from Antarctica in 2003. It gathered data for 14 days in three frequency bands, centered at 145, 245, and 345 GHz. Here we restrict ourselves to the 145 GHz data, which are most sensitive to CMB fluctuations. These have been obtained with polarization-sensitive bolometers (PSBs). The analysis of the data set has produced high-quality maps of the southern sky that have been conveniently divided in three regions: a “deep” (in terms of integration time) survey region ($\sim 90$ deg$^2$) and a “shallow” survey region ($\sim 750$ deg$^2$), both at high Galactic latitudes, as well as a region of $\sim 300$ deg$^2$ across the Galactic plane. The deep region is completely embedded in the shallow region. Here we only consider a subset of the data that contains all of the deep region and part of the shallow, for a total of 693 deg$^2$ ($1.7\%$ of the sky). The mask we use is square, 26 deg on a side, centered at about R.A. = 82.6$^\circ$ and decl. = $-44.2^\circ$, and excludes all detected point sources in the field. This region has been selected in view of high-sensitivity CMB observation with low foreground contamination and was observed with a highly connected scanning strategy to keep systematics under control. We use the $T$ map data reduced jointly from eight PSBs at 145 GHz (Masi et al. 2006). In this region, the signal rms on $3.4^\circ$ pixels is $\sim 90$ $\mu$K and instrumental noise has an rms of $\sim 20$ $\mu$K in the deep region and $\sim 90$ $\mu$K in the shallow region. In harmonic space, binned estimates of the CMB angular power spectrum retain signal-to-noise ratios $>1$ well beyond $l \sim 1000$. One may compare these figures with WMAP: in the 3 year release, the WMAP combined sensitivity in the region observed by B03 is $\sim 100$ $\mu$K on $3.4^\circ$ pixels, close to WMAP’s mean pixel error. However, WMAP’s beams are broader than B03, so its $l$-space error is $\sim 5$ times larger than B03 at $l \sim 1000$. On the other hand, B03 was not devised to measure multipoles at $l \leq 50$. In this sense, our NG analysis probes angular scales complementary to those constrained by WMAP.

While we do not consider here the Stokes $Q$ and $U$ polarization maps, our $T$ map has been marginalized with respect to linear polarization. For a description of the instrument and the measured $T$ and polarization maps, see Masi et al. (2006), and for the CMB TT, TE, and EE power spectra, see Jones et al. (2006), Placentini et al. (2006), and Montroy et al. (2006).

To assess the robustness of our tests of Gaussianity we use a set of simulated MC maps that mimic the B03 data. To produce these, we follow the same steps performed when analyzing real data. The Gaussian CMB sky signal is simulated from the power spectra that best fit the B03 data (MacTavish et al. 2006). This signal is smoothed according to the measured beam and synthesized into a pixelized sky map, using HEALPix routines (Górski et al. 2005). Using the B03 scanning strategy, the signal map is projected onto eight time streams, one for each 145 GHz detector. Noise-only time streams are also produced, as Gaussian realizations of each detector’s noise power spectral density, which are estimated from the data accounting for cross talk among detectors. The time lines are reduced with the ROMA map-making code (Natoli et al. 2001; de Gasperis et al. 2003). The power spectrum of the NG maps is identical to that of the Gaussian CMB simulations.

3. TESTS OF GAUSSIANITY AND CONSTRAINTS ON $f_{\text{NL}}$

Working at 6.8’ HEALPix resolution ($N_{\text{side}} = 512$), we first compute the normalized skewness $S_3$ and kurtosis $S_4$ of our pixelized field $T_i$. These are obtained from the variance $\sigma^2 = 1/(N-1) \sum (T_i - \langle T \rangle)^2$ and from the third and the fourth moments $\mu_3 = \sum (T_i - \langle T \rangle) \langle T \rangle^2/N$ and $\mu_4 = \sum (T_i - \langle T \rangle)^4/N$, where $N$ is the total number of pixels of the map and $\langle T \rangle = \sum T_i/N$ its mean. We have $S_3 = \mu_3/\sigma^3$, $S_4 = \mu_4/\sigma^4 - 3$. From the data we get $S_3 \sim -0.016$ and $S_4 \sim 0.096$. These values are plotted in Figure 1 as a vertical line and compared to the empirical distribution as derived from the MC (signal plus noise) maps. From the latter we compute the probabilities $P(S_3^{\text{sim}} > S_3^{\text{data}}) = 58\%$ and $P(S_4^{\text{sim}} > S_4^{\text{data}}) = 77\%$. Hence, for these tests the data are compatible with the Gaussian hypothesis. The same tests are repeated after having degraded the map to $13.6'$, finding similar results.

To analyze the map with MFs (Gott et al. 1990), we consider

![Image](image-url)
the excursion sets $Q$ defined as the map’s subsets exceeding a given threshold $v$: $Q(v) = \{ T ; (T - \langle T \rangle)/\sigma > v \}$. The three MFs measure the total area of the surviving regions of the map ($M_f$), their total contour length ($M_c$), and the genus of the distribution, which is related to the difference between the number of “hot” and “cold” regions ($M_g$). For a Gaussian field the expectation values of the functionals depend on a single parameter $\tau$. $\langle M_f(v) \rangle = \frac{1}{2} [1 - \text{erf}(v/\sqrt{2})]$, $\langle M_c(v) \rangle = (7/8) \exp(-v^2/2)$, $\langle M_g(v) \rangle = \tau/(8\pi^3/3) \exp(-v^2/2)$. In the case of a pure CMB signal (no noise), $\tau$ is given by $\tau = \frac{1}{2} \sum_{k,l} (2l+1) \delta(k) C_l$ (Schmalzing & Górski 1998; Winitzki & Kosowsky 1998).

Here $M_f$ ($M_c$) is any of the three MFs obtained from the data (simulations), $\langle \cdot \rangle$ is the mean value over MC realizations, and $C_{ij} = \langle (M_i - \langle M_i \rangle)(M_j - \langle M_j \rangle) \rangle$ is a covariance matrix, estimated from an independent set of $\sim 1000$ Gaussian maps. In the top left, top right, and bottom left panels of Figure 3 we show $\chi^2$ for each MF (vertical line), plotted along with the empirical distribution sampled via MC. We can define a “joint” estimator by grouping the $M_i$’s in a single, 27 element data vector $M_{ij} = [M_{f0}, M_{c0}, M_{g0}]$ and defining a $\chi^2$ as a trivial extension of equation (2). It is important that the covariance matrix built for the joint estimator correctly account for correlations among different functionals. However, we have found that in order to pin down to a stable level these off-block couplings, one requires a number realizations significantly higher than the $\sim 1000$ used throughout our analysis. The latter number cannot be realistically increased to desired level, because the GLS map-making procedure is a demanding computational task, even for the supercomputers we have used. Fortunately, we have found that using white (instead of correlated) noise to estimate the covariance matrix has a negligible impact on the analysis. This finding can indeed be justified a posteriori, because the GLS map-making procedure is very effective in suppressing noise correlations, which contribute very weakly to the estimator’s final covariance. The joint $\chi^2$ of the data is displayed as the bottom right panel in Figure 3, along with the MC empirical distribution. The probability $P(\chi^2 > \chi^2_{\text{obs}})$ that a Gaussian map has a larger $\chi^2$ than the B03 map is 76% for $M_f$, 83% for $M_c$, 76% for $M_g$, and 67% for the “joint” estimator. The values are fully consistent with the Gaussian hypothesis. We conclude that our pixel-space analysis does not detect any sign of NG behavior in the B03 data.

The problem is to use all of the statistical information contained in the $\chi^2$ distribution for the MFs. 

We now want to constrain the quadrupolar coupling parameter $f_{NL}$ defined in equation (1). To this purpose we simulate NG CMB realizations in the following way: first, we generate the Gaussian and NG part of the primordial potential defined by equation (1); then we convolve them with CMB first-order radiation transfer functions to get the final CMB sky. In this way we produce 200 G (Gaussian) maps and 200 NG counterparts (each G map has a uniquely defined NG counterpart), so that for a given $f_{NL}$ our (signal only) map is $G + f_{NL} \times \text{NG}$. By adding noise maps, we can define the MF estimators as discussed above, with the difference that the are now functions of $f_{NL}$: $fB_{NL}(f_{NL}) = M_f - (M_f(f_{NL}))$ (we only consider the “joint” estimator in what follows). Consequently, we now define the data $\chi^2$ as $\chi^2_{\text{obs}}(f_{NL}) = fB_{NL}(f_{NL})^T C fB_{NL}(f_{NL})$. While in principle the covariance of the $M_i$’s is a function of $f_{NL}$, this dependence is expected to be weak and is usually neglected (Komatsu et al. 2003). We have tested for this explicitly by using our NG simulations. We plot $\chi^2_{\text{obs}}$ as a function of $f_{NL}$ in Figure 4 (left panel). Goodness-of-fit analysis yields $-300 < f_{NL} < 650$ at 68% CL and $-800 < f_{NL} < 1050$ at 95%, with a best-fit value ($\chi^2_{\text{obs}}$ at its minimum) of $f_{NL} = 200$. In order to cross-check this result, we estimate a “frequentist” confidence interval for $f_{NL}$ by sampling the empirical distribution of the $M_f$-based $\chi^2$, computed for $f_{NL} = 200$. The resulting histogram is displayed in the right panel of Figure 4. This analysis shows that we should expect to constrain $|f_{NL}| \leq 1000$ at 95%, thus suggesting that our limits derived through goodness-of-fit analysis are consistent. One may consider what limits on $f_{NL}$ would be derived if we use, in place of MFs, the map’s skewness and kurtosis defined above as elements of a two-dimensional data vector. We thus repeated our goodness-of-fit analysis using
these statistics and found weaker limits: \(-950 < f_{NL} < 1350\) at 68% CL (\(-1400 < f_{NL} < 1800\) at 95% CL). Even so, it is quite remarkable that a crude one-point pixel statistic degrades the final sensitivity only by a factor \(\sim 2\). Of course, in order to find “optimal” constraints one has to implement a more sophisticated analysis.

The B03 data set is not expected to be noise dominated at the 6.8′ resolution employed in the analysis above. To show that this is the case, we repeated all of our procedures for 13.6′, finding similar (although slightly weaker) constraints.

4. CONCLUSIONS

We analyzed the B03 145 GHz T map in search of NG signals. We worked in pixel space at 6.8′ and 13.6′ HEALPix resolution. We computed the skewness and kurtosis of the map, as well as its three MFs. We compared these estimates against a set of simulated Gaussian maps that have been reduced using the same analysis pipeline as real data, finding no evidence of NG behavior. To quantify the latter statement, we define goodness-of-fit statistics jointly based on all three MFs, showing that the probability for a Gaussian simulation to have a \(\chi^2\) larger than the data is \(\sim 67\%\). Assuming a model for primordial fluctuations that predicts a quadratic perturbation to the gravitational potential, we set limits on the nonlinear coupling parameter \(f_{NL}\) as \(-800 < f_{NL} < 1050\) at 95% CL (\(-300 < f_{NL} < 650\) at 68% CL). These limits may be regarded as complementary to the constraints set by WMAP in view of the better signal-to-noise ratio at high resolution in the B03 field.

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