THE VECTOR VORTEX CORONAGRAPH: LABORATORY RESULTS AND FIRST LIGHT AT PALOMAR OBSERVATORY

D. MAWET1,4, E. SERABYN1, K. LIEVER1, R. BURRUS1, J. HICKEY2, AND D. SHEMO3
1 Jet Propulsion Laboratory, California Institute of Technology, Pasadena, CA 91109, USA; Dimitri.Mawet@jpl.nasa.gov
2 Palomar Observatory, California Institute of Technology, P.O. Box 200, Palomar Mountain, CA 92060, USA
3 JDS Uniphase Corporation, 2789 Northpoint Parkway, Santa Rosa, CA 95407, USA

ABSTRACT

High-contrast coronagraphy will be needed to image and characterize faint extrasolar planetary systems. Coronagraphy is a rapidly evolving field, and many enhanced alternatives to the classical Lyot coronagraph have been proposed in the past 10 years. Here, we discuss the operation of the vector vortex coronagraph, which is one of the most efficient possible coronagraphs. We first present recent laboratory results and then first light observations at the Palomar observatory. Our near-infrared H-band (centered at ~1.65 μm) and K-band (centered at ~2.2 μm) vector vortex devices demonstrated excellent contrast results in the lab, down to ~10^{-6} at an angular separation of ~3λ/d. On sky, we detected a brown dwarf companion 3000 times fainter than its host star (HR 7672) in the Ks band (centered at ~2.15 μm), at an angular separation of ~2.5λ/d. Current and next-generation high-contrast instruments can directly benefit from the demonstrated capabilities of such a vector vortex: simplicity, small inner working angle, high optical throughput (>90%), and maximal off-axis discovery space.

Key words: brown dwarfs – instrumentation: adaptive optics – instrumentation: high angular resolution – techniques: high angular resolution

Online-only material: color figures

1. INTRODUCTION

High-contrast imaging of extrasolar planetary systems represents one of the most challenging technological goals of modern observational astrophysics. The challenge is twofold: the small angular separation between exoplanets and their host stars requires a high angular resolution (e.g., 100 mas for a 1 AU distance at 10 pc), and the brightness contrast between them is tremendous, ranging from 10^{-3} for very close hot giant planets up to 10^{-10} for Earth-like planets. After the first resolved observation of a bound planetary mass object was obtained by Chauvin et al. (2005), several exoplanets have recently been imaged. The majority of these observations were eased by the fairly large distance of the companions from their host stars (0.5–15”) and the moderate star/planet contrast in the near-infrared (Δm ≈ 5–12) due to the relatively high thermal emission of the young forming planets (at ~1000 K). The best characterizations obtained so far have relied on accurate multi-wavelength broadband photometry in the visible and near-infrared, and on multi-epoch astrometry. Keplarian motion was unambiguously detected for the most recent discoveries of Kalas et al. (2008) around Fomalhaut and Marois et al. (2008) around HR 8799.

Direct imaging is clearly a self-consistent and complete characterization method as it provides a straightforward means to constrain orbital parameters and perform spectro-photometric measurements in the most varied orbital configurations, at almost any point in time. However, the few exoplanets and the majority of circumstellar disks imaged so far all required high contrast (Absil & Mawet 2010; Oppenhime & Hinkley 2009; Duchêne 2008; Wyatt 2008). Image dynamic ranges can be increased with a coronagraph, and in a comparison of coronagraphs, Guyon et al. (2006) identified two as closest to theoretically ideal in terms of “useful throughput”: the phase-induced amplitude apodization coronagraph and the optical vortex phase-mask coronagraph (Mawet et al. 2005; Foo et al. 2005; Jenkins 2008; Swartzlander et al. 2008). The optical vortex is truly close to ideal, as it has a small inner working angle (IWA), high throughput, and a completely clear off-axis discovery space. We have therefore been pursuing an especially promising version of the vortex coronagraph, the vector vortex coronagraph (VVC).

Here, we present our most recent laboratory results with VVC masks based on birefringent liquid crystal polymers (LCPs). These tests were carried out in the near-infrared, as the masks considered here are intended for use behind ground-based extreme adaptive optics systems (ExAO). We also present the results and consider the implication of the first use of our VVC on the sky with the Palomar ExAO-level well-corrected subaperture (WCS; Serabyn et al. 2007, 2009).

2. THE VECTOR VORTEX CORONAGRAPH

An optical vortex is generated by a phase screw of the form eiθ, with θ being the azimuthal coordinate. There are two kinds of optical vortex known: the scalar optical vortex, implemented by a structural helix (Foo et al. 2005; Swartzlander et al. 2008) providing a scalar optical phase delay which applies to the two orthogonal polarization components of natural light identically, and the vectorial vortex, implemented by a rotationally symmetric halfwave plate (HWP), providing a “geometrical” phase shift (Pancharatnam 1956; Berry 1987) that applies opposite phase screws to the two orthogonal circular polarization states (Figure 1).

The detailed theory of the VVC is provided in Mawet et al. (2005, 2009a), and so here we merely provide a quick schematic overview. In the vector vortex, for a linearly polarized input field (or for natural light projected onto a linear basis), the
Figure 1. VVC azimuthal phase ramp. Panel (a): rotationally symmetric HWP with an optical axis orientation that rotates about the center (dashed lines perpendicular to the circumference). The net effect of an HWP on an impinging linear polarization is to rotate it by \( -2 \times \alpha \), where \( \alpha \) is the angle between the incoming polarization direction and the fast optical axis. An incoming horizontal polarization (blue arrow) is transformed by the vector vortex so that it spins around its center twice as fast as the azimuthal coordinate \( \theta \) (red arrows). Panel (b): for circular polarization, the output field rotation is strictly equivalent to a phase delay (the starting angle 0 is rotated; therefore phase-shifted). The angle of local rotation of the polarization vector corresponds to a “geometrical” phase: upon a complete rotation about the center of the rotationally symmetric HWP, it has undergone a total \( 2 \times 2\pi \) phase ramp, which corresponds to the definition of an optical vortex of topological charge 2 (panel (c)). Upon propagation from the focal plane to the subsequent pupil plane, the Fourier transform (FT) of the product of the PSF by the azimuthal phase ramp sends the light outside the original pupil area (panel (d)). (A color version of this figure is available in the online journal.)

A rotationally symmetric HWP rotates the polarization vector as in Figure 1(a). The definition of circular polarization is just a linear polarization rotating at the angular frequency \( \omega \) (equal to the pulsation of the electromagnetic field), so that a rotation \( \phi = 2\theta \) of the polarization vector is strictly equivalent to a phase delay (Figure 1(b)). If, at any given point in space, the polarization vector is rotated as shown in Figure 1(a), it implies for the circular polarization (Figure 1(b)) that it has acquired a geometric phase ramp \( e^{i\phi} = e^{i2\theta} \); such as represented in Figure 1(c). Thus, \( \phi \) represents both an angle and a phase—hence the term “geometrical” phase.

The factor 2 before the azimuthal coordinate in \( e^{i2\theta} \) is called the “topological charge” \( l \). It determines the polarization spin rate and the subsequent height of the phase ramp after a full \( 2\pi \) rotation. In this particular example, we considered a vortex of topological charge 2, meaning that, upon a complete rotation about the center, it has undergone a total \( 2 \times 2\pi \) phase shift. Note that this azimuthal phase ramp is generated without a corresponding structural helix. When the point-spread function (PSF) at the focal plane of a telescope is centered on such a vector vortex, for non-zero even values of \( l \), the light, upon propagation to a subsequent pupil plane, then appears entirely outside the original pupil area, where it can be rejected by a Lyot stop of the same size or slightly smaller than the entrance pupil (Figure 1(d); Mawet et al. 2005).

3. LCP TECHNOLOGY AND CURRENT DEVICES

We have been developing rotationally symmetric HWP devices based on LCPS, which combine the birefringent properties of liquid crystals with the excellent mechanical properties of polymers. The orientation of a LCP is achieved through photoalignment. Once aligned and cured, the polymer reaches a very stable solid state. The detailed proprietary fabrication process of our manufacturer JDSU has already been described in Mawet et al. (2009a). In that paper, we presented the theory of operation and obtained encouraging results, experimentally validating the VVC concept and showing good agreement with an empirical model of the imperfections. The first sample was mainly limited by the size (diameter \( \sim 100 \mu m \)) of the central so-called region of disorientation, where the LCP loses the appropriate orientation close to the vortex axis, where most of the energy of the Airy diffraction pattern of the central stellar source would land. Our goal for the next generation of devices was thus to significantly decrease the central region of disorientation, and to cover the reduced residual zone by an opaque aluminum mask so as to prevent leakage through the center of the mask. Numerical simulations in Mawet et al. (2009a) promised dramatic improvements in contrast if these steps were taken. The reduction was to be enabled by a redesign of the mechanical manufacturing apparatus described in Mawet et al. (2009a).

The new topological charge 2 \( H \)- and \( K \)-band devices achieved our manufacturing goals. The size of the central region of disorientation was reduced to a diameter of \( \sim 25 \mu m \), a four-fold improvement upon the first vortex device. We chose to cover the remaining residual zone by an aluminum spot 25 \( \mu m \) in diameter (Figure 2, right). Given any beam \( F \) number \( \geq 10 \) (\( F = f/d \), with \( f \) the focal length and \( d \) the diameter of the beam as defined by the entrance pupil), the radius of the spot is comparable to or smaller than the beam width \( F \lambda \) (where \( \lambda \) is the wavelength of the incident light), so the Al spot does not affect the IWA at all. Note that classical and apodized Lyot coronagraphs use much larger central opaque spots, many \( F \lambda \) across. The polarization properties of the new devices were measured on an Axoscan Mueller Matrix Spectro-Polarimeter. These measurements allowed us to calculate the central wavelengths (where the retardance is \( \pi \) radians) of the \( H \) and \( K \) devices to be 1642.5 nm and 2225.2 nm, respectively, reasonably well centered on the \( H \) and \( K \) bands. Given the birefringence dispersion of the LCP material, one can calculate the standard deviation of the retardance over the filter fractional bandpass, \( \epsilon \). An estimate of the total attenuation \( A \) (integrated over the whole field of view) can then be calculated from \( \epsilon \) (Mawet et al. 2009a); \( A = \epsilon^2/4 \). Table 1 summarizes our total attenuation predictions and the peak-to-peak attenuation measurements presented in Section 4. Indeed, a coronagraph’s ability to suppress the starlight is conveniently quantified and measured by the peak-to-peak attenuation, which is the ratio of the maximum in the post-coronagraphic stellar residual image to the maximum (on-axis)
intensity in the direct stellar image. This metric is especially meaningful on sky, since its inverse directly represents the allowable gain in integration time for individual unsaturated exposures. The peak-to-peak attenuation is close to but can be slightly smaller than the total attenuation (integrated over the whole image), owing to the spreading of the residual PSF in an aberrated system.

4. LAB SETUP AND RESULTS

We tested our new devices at the Jet Propulsion Laboratory (JPL) on our transmissive infrared coronagraphic testbench (Mawet et al. 2009a). To simulate starlight, we used a supercontinuum white-light source (Fianium SC450) coupled to standard telecom monomode fibers (SMF-28e) to provide a clean spatially filtered wave front as input to the coronagraph. The testbench uses standard telecom monomode fibers (SMF-28e) to provide a clean spatially filtered wave front as input to the coronagraph. The entrance pupil is defined by an iris diaphragm. To filter the pupil (Lyot) plane, we used another diaphragm with a diameter set to be smaller than the entrance pupil by about 10% (providing an optical throughput of ∼81%). All images were recorded with a LN2 cooled 12 bit Merlin InSb camera. The cold filter in the camera is a Spectrogon lowpass filter (cutoff ∼2.6 μm). We oversampled the images relative to $F\lambda$ to allow a proper peak-to-peak attenuation measurement. We adopted an $F$ number of 40, similar to that planned for future telescope instruments such as SPHERE (Beuzit et al. 2007). We first built up a coronographic image with the light source image centered on the coronagraphic spot by adding 100 individual exposure of 16 ms. The number of frames and the individual exposure time are chosen to balance signal with background and readout noise. The next step was to record direct non-coronagraphic images (the mask was translated off the focal plane spot) in the same conditions. All images had a median dark frame subtracted from them.

Results in terms of azimuthally averaged coronagraphic profiles (radial cut in the non-attenuated and attenuated PSFs) are shown for the $K$-band mask in Figure 3. The peak-to-peak attenuation is ∼4 × 10$^{-3}$ in the 15% bandwidth and ∼2.5 × 10$^{-4}$ in the reduced 5% equivalent filter. The contrast near 3λ/d angular separation is 3 × 10$^{-5}$ ± 0.05 × 10$^{-5}$ for 15% and 2 × 10$^{-6}$ ± 0.5 × 10$^{-6}$ for 5% bandwidth. In the $H$ band, we obtained similar results (∼5 × 10$^{-3}$ peak-to-peak attenuation). Our speckle noise level on the bench is ∼5 × 10$^{-7}$, which can be considered as the error bar on all our contrast measurements. Our laboratory results are consistent with our model of the device, taking into account the residual chromaticity of the LCP material (Section 3), as well as the diffraction effect of the aluminum spot, which is negligible here (Table 1). These excellent results demonstrate that these devices are ready for use on a telescope, where the atmosphere will likely be the main limitation.

5. ON-SKY RESULTS

We next installed our $K$-band VVC mask in the cryogenic infrared camera PHARO (Hayward et al. 2001). The first tests on the 5 m Hale telescope were performed with the WCS on
2009 May 30. The Palomar WCS has been described earlier (Serabyn et al. 2007). In short, an unvignetted, off-axis WCS is provided by means of a set of relay optics upstream of the Palomar AO system (PALAO). The off-axis relay magnifies and shifts an off-axis subaperture pupil onto the deformable mirror, yielding roughly 10 cm actuator spacing in the selected pupil. This configuration reduces the effective aperture size from 5 m to 1.5 m, but allows operation in the ExAO regime. While providing unprecedented image quality on a ground-based telescope (Strehl ratios, S \geq 90\%), this off-axis setup is also ideal for use with phase-mask coronagraphs because these coronagraphs work best with unobscured apertures, and their very small IWA partly compensates for the loss in resolution due to the smaller aperture.

We first conducted contrast measurements with PALAO and PHARO attached to the Cassegrain focus of the 5 m Hale telescope and with PALAO locked on the internal white-light source. The Strehl ratio of this instrumental configuration is S \approx 99\%. Using PHARO’s 14\% wide K\_ filter, we measured a peak-to-peak attenuation \sim 3 \times 10^{-3}, comparable to our lab measurement (Figure 3 and Table 1).

We next pointed the telescope at the star HR 7672 (K = 4.4, V = 5.8), which possesses a brown dwarf companion (Liu et al. 2002). The companion has a small angular offset and high contrast ratio, r = 790 \pm 5 mas and \Delta K = 8.62 \pm 0.07 mag (\sim 3 \times 10^{-4}), making it the ideal test case target. We took a series of twenty 4 s exposures with the VVC centered on the target to within a few mas. We immediately followed this sequence by the same series of exposures on a reference star, with the same V – K color and brightness as the target. Matching V magnitudes is important to ensure a similar AO correction. K magnitudes also need to be matched to ensure a proper signal-to-noise ratio match for the quasi-static speckle subtraction. Also, to reduce orientation and flexure-induced variations as much as possible, the calibration star was chosen and observed at the same hour angle and elevation as the target star.

Figure 4 (left) shows our image of HR 7672 after post-processing with the locally optimized combination of images (LOCI) algorithm (Lafrenière et al. 2007), which is one of the most effective speckle subtraction techniques. The vortex coronagraph, owing to its absence of features anywhere in the field of view, allowed us to make an even cleaner detection (in only 80 s) than the four-quadrant phase mask (Serabyn et al. 2009), even in the presence of poor seeing. Note that for the WCS, the companion lies at only \sim 2.5\_ / d.

Figure 4 (right) shows the radial cut in the unattenuated PSF (black), and contrast curves, starting with the raw on-sky contrast (continuous blue curve), and its associated 3\sigma photon noise for an 80 s integration time (dashed blue line). The dashed green curve represents the detection level of the final image after LOCI processing, which effectively brings us very close to the photon noise limit.

Note that the peak-to-peak attenuation (\sim 4 \times 10^{-2}) is an order of magnitude worse than the instrumental limit measured on the internal AO source (\sim 3 \times 10^{-3}). This difference can be explained by three major observational limitations.

1. The seeing was an average of 1\arcsec, and S was \sim 85\%. S allows us to roughly estimate the attenuation by substituting the residual wavefront phase variance \sigma^2 from Marechal’s approximation S = e^{-\sigma^2}, in the attenuation formula of Section 3 (Boccaletti et al. 2004), yielding A = - \ln S / 4 \approx 4 \times 10^{-2}, which is consistent with the measured peak-to-peak attenuation.

2. Residual pointing offsets affected the consistency of the centering of the star on the mask.

3. On the sky, the vortex was slightly out of focus. The first limitation is contingent upon the atmospheric turbulence conditions, i.e., the seeing and the turbulence coherence time. The last two are instrumental systematics that can be corrected to yield an improvement in raw contrast. Our 3\sigma projected instrumental contrast limit (for a K = 4.4 mag star in 1800 s)
extrapolated from the white-light source contrast measurement is shown in Figure 4 (right, dashed red curve). This limit can be reached in practice under more favorable seeing conditions (<1″).

6. CONCLUSIONS

We have demonstrated that a vector vortex is an extremely powerful and yet simple coronagraph. The discovery space of the VVC is maximal, with a small IWA (∼1.5λ/d) and a theoretical optical throughput larger than 90%. Thanks to ground-breaking advances in LCP technology, a set of usable high performance VVC devices was produced in a very short time. We have demonstrated laboratory contrasts on the order of a few 10^{-5} at an angular separation of ∼3λ/d over 15% bandwidth. Our best result in a 5% equivalent bandwidth is ∼2 × 10^{-6} at 3λ/d. On sky, we easily detected a known brown dwarf at ∼2.5λ/d (for the WCS), with a contrast ratio of ∼3 × 10^{-3}. Our current generation of near-infrared devices is in fact already compliant with the specifications of current and next-generation ground-based instruments for exoplanet imaging and characterization such as SPHERE (Beuzit et al. 2007), GPI (Macintosh et al. 2007), Palm-3000 (Dekany et al. 2006), and HiCIAO (Tamura et al. 2006).

The LCP technology also applies at the optical wavelengths, and we have been pursuing that route as well. The tests of our initial visible device showed good contrast results on the High Contrast Imaging Testbed (Trauger & Traub 2007). Using wave front control and a topological charge 4 VVC centered at 800 nm, a dark hole with ∼10^{-7} raw contrast was created in a 10% bandpass (Mawet et al. 2009b). Our further development of improved visible devices is ultimately aimed at providing a close to ideal coronagraph for exoplanet space missions such as the Actively Corrected Coronagraph for Exoplanet Space Studies (Trauger et al. 2008), paving the way toward the Terrestrial Planet Finder Coronagraph (Traub et al. 2006).

This work was carried out at the Jet Propulsion Laboratory (JPL), California Institute of Technology (Caltech), under contract with NASA. D.M. is supported by an appointment to the NASA Postdoctoral Program at JPL, Caltech, administered by Oak Ridge Associated Universities through a contract with NASA. The data presented in this paper are based on observations obtained at the Hale Telescope, Palomar Observatory, as part of a continuing collaboration between Caltech, NASA/JPL, and Cornell University.

REFERENCES

Absil, O., & Mawet, D. 2010, A&AR, in press
Berry, M. 1987, J. Mod. Opt., 34, 1401
Beuzit, J.-L., et al. 2007, in Proc. In the Spirit of Bernard Lyot: The Direct Detection of Planets and Circumstellar Disks in the 21st Century, ed. P. Kalas (Berkeley, CA: Univ. California Press)
Boccaletti, A., Riaud, P., Baudou, P., Baudrand, J., Rouand, D., Gratadour, D., Lacombe, F., & Lagrange, A.-M. 2004, PASP, 116, 1061
Chauvin, G., Lagrange, A.-M., Dumas, C., Zuckerman, B., Mouillet, D., Song, I., Beuzit, J.-L., & Lowrance, P. 2005, A&A, 438, L25
Dekany, R., Bouchet, A., Britton, M., Velur, V., Troy, M., Shelton, J. C., & Roberts, J. 2006, Proc. SPIE, 6272, 62720G
Duchêne, G. 2008, New Astron. Rev., 52, 117
Foo, G., Palacios, D. M., & Swartzlander, Jr., G. A. 2005, Opt. Lett., 30, 3308
Guyon, O., Pluzhnik, E. A., Kuchner, M. J., Collins, B., & Ridgway, S. T. 2006, ApJS, 167, 81
Hayward, T. L., Brandl, B., Pirger, B., Blacken, C., Gull, G. E., Schoenwald, J., & Houck, J. R. 2001, PASP, 113, 105
Jenkins, C. 2008, MNRRAS, 384, 515
Kalas, P., et al. 2008, Science, 322, 1345
Lafrenière, D., Marois, C., Doyon, R., Nadeau, D., & Artigau, É. 2007, ApJ, 660, 770
Liu, M. C., Fischer, D. A., Graham, J. R., Lloyd, J. P., Marcy, G. W., & Butler, R. P. 2002, ApJ, 571, 519
Macintosh, B., et al. 2007, C. R. Phys., 8, 365
Marois, C., Macintosh, B., Barman, T., Zuckerman, B., Song, I., Patience, J., Lafrenière, D., & Doyon, R. 2008, Science, 322, 1348
Mawet, D., Riaud, P., Absil, O., & Surdej, J. 2005, ApJ, 633, 1191
Mawet, D., Serabyn, E., Liewer, K., Hanot, C., McElidowney, S., Shemo, D., & O’Brien, N. 2009a, Opt. Express, 17, 1902
Mawet, D., Trauger, J., Serabyn, E., Moody, D., Liewer, K., Krist, J., Shemo, D., & O’Brien, N. 2009b, Proc. SPIE, 7440, 74400X
Oppenheimer, B. R., & Hinkley, S. 2009, ARA&A, 47, 253
Pancharaman, S. 1956, Proc. Indian Acad. Sci., 44, 247
Serabyn, E., Mawet, D., Bloemhof, E., Huguenauner, P., Mennesson, B., Wallace, K., & Hickey, J. 2009, ApJ, 696, 40
Serabyn, E., Wallace, K., Troy, M., Mennesson, B., Huguenauner, P., Gappinger, R., & Burress, R. 2007, ApJ, 658, 1386
Swartzlander, G. A., Jr., Ford, E. L., Abdul-Malik, R. S., Close, L. M., Peters, M. A., Palacios, D. M., & Wilson, D. W. 2008, Opt. Express, 16, 10200
Tamura, M., et al. 2006, Proc. SPIE, 6269, 62690V
Traub, W. A., et al. 2006, Proc. SPIE, 6268, 62680T-1
Trauger, J. T., & Traub, W. A. 2007, Nature, 446, 771
Trauger, J., et al. 2008, Proc. SPIE, 7010, 701029-1
Wyatt, M. C. 2008, ARA&A, 46, 339