Hubble Space Telescope imaging of the extremely metal-poor globular cluster EXT8 in Messier 31 *
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Søren S. Larsen¹, Aaron J. Romanowsky²,³, and Jean P. Brodie³,⁴

¹ Department of Astrophysics/IMAPP, Radboud University, PO Box 9010, 6500 GL Nijmegen, The Netherlands
e-mail: s.larsen@astro.ru.nl
² Department of Physics & Astronomy, One Washington Square, San José State University, San Jose, CA 95192, USA
³ Centre for Astrophysics and Supercomputing, Swinburne University of Technology, Hawthorn, VIC 3122, Australia
⁴ University of California Observatories, 1156 High Street, Santa Cruz, CA 95064, USA

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ABSTRACT

We recently found the globular cluster (GC) EXT8 in M 31 to have an extremely low metallicity of [Fe/H] = −2.91 ± 0.04 using high-resolution spectroscopy. Here we present a colour–magnitude diagram (CMD) for EXT8, obtained with the Wide Field Camera 3 on board the Hubble Space Telescope. Compared with the CMDs of metal-poor Galactic GCs, we find that the upper red giant branch (RGB) of EXT8 is ~0.03 mag bluer in M_F606W−M_F814W and slightly steeper, as expected from the low spectroscopic metallicity. The observed colour spread on the upper RGB is consistent with being caused entirely by the measurement uncertainties, and we place an upper limit of σ_F606W−F814W = 0.015 mag on any intrinsic colour spread. The corresponding metallicity spread can be up to σ[Fe/H] ~ 0.2 dex or >0.7 dex, depending on the isochrone library adopted. The horizontal branch is located mostly on the blue side of the instability strip and has a tail extending to at least M_F606W = +3, as in the Galactic GC M 15. We identify two candidate RR Lyrae variables and several ultraviolet-luminous post-horizontal-branch and/or post-asymptotic-giant-branch star candidates, including one very bright (M_F390X = −3.2) source near the centre of EXT8. The surface brightness of EXT8 out to a radius of 25′′ is well fitted by a Wilson-type profile with an ellipticity of £ = 0.20, a semi-major axis core radius of 0.′25, and a central surface brightness of mean [M_F606W] = 15.2 mag arcsec−2, with no evidence of extra-tidal structure. Overall, EXT8 has properties consistent with it being a ‘normal’, but very metal-poor, GC, and its combination of relatively high mass and very low metallicity thus remains challenging to explain in the context of GC formation theories operating within the hierarchical galaxy assembly paradigm.

Key words. globular clusters: individual: RBC EXT8 – Hertzsprung-Russell and C-M diagrams – stars: horizontal-branch – stars: AGB and post-AGB

1. Introduction

In his investigation of globular cluster (GC) radial velocities, Mayall (1946) noted that the integrated-light spectral types of GCs ranged from fairly early (A5) to roughly solar (G5). During the following decade, several authors linked the extremely weak metal lines in the spectra of GCs such as M 15 and M 92 to a deficiency in the abundances of the corresponding elements (Baum 1952; Morgan 1956; Baade 1958), and the first quantitative spectral analyses established that some GCs have metal/hydrogen ratios of less than 1% of the solar value (Helfer et al. 1959; Kinman 1959). Measurements of chemical abundances and differences in the spatial distributions of metal-poor and metal-rich GCs were discussed in the context of the then-nascent fields of nucleosynthesis and Galactic chemical evolution (Burbridge et al. 1957; Helfer et al. 1959; Morgan 1959), leading up to the seminal work of Searle & Zinn (1978). Around the same time, classical investigations of GC colour–magnitude diagrams (CMDs) revealed correlations between metallicity and CMD characteristics such as the slope of the red giant branch (RGB) and the morphology of the horizontal branch (HB) (Arp 1955; Sandage & Wallerstein 1960; Sandage & Smith 1966). These results played an important role in informing early stellar evolutionary models (Hoyle & Schwarzschild 1955; Kippenhahn et al. 1958; Demarque & Geisler 1963). It was thus established early on that GCs can provide important insights into both stellar and Galactic astrophysics.

As more data became available, the relatively low incidence of clusters with metallicities near [Fe/H] = −2.5 became evident (Bond 1981). By comparison with simple models for Galactic chemical evolution or with metallicity distributions of halo field stars, it has been estimated that there should be a handful of GCs with [Fe/H] < −2.5 in the Milky Way, while none is observed (Carney et al. 1996; Simpson 2018; Beasley et al. 2019; Youakim et al. 2020). This has led to the notion of a ‘metallicity floor’ for GCs near [Fe/H] ≈ −2.5. It has been suggested that such a metallicity floor may be a consequence of the hierarchical nature of galaxy assembly, combined with the galaxy mass–metallicity relation (Harris et al. 2006; Choksi et al. 2018; Usher et al. 2018; Kroupa et al. 2019). In this scenario, dwarf (proto-) galaxies with metallicities less than [Fe/H] = −2.5 that contributed to the build-up of galactic halos would be insufficiently massive to form massive GCs that could survive for a Hubble time. These ideas appeared to be supported...
by the identification of the Phoenix stream in the Milky Way halo as the possible remnant of a low-mass GC with [Fe/H] = −2.7 (Wan et al. 2020).

From an integrated-light spectrum obtained with the HIRES spectrograph on the Keck I telescope, the GC EXT8 in Messier 31 (M 31) was recently shown to have a metallicity of [Fe/H] = −2.91 ± 0.04, well below the metallicity floor (Larsen et al. 2020). Unlike the Phoenix stream progenitor, EXT8 is fairly massive, with an estimated dynamical mass of (1.14 ± 0.16) × 10^6 M⊙, and thus represents a challenge to the notion that metal-poor, massive GCs could not have formed in the early Universe. The cluster was first mentioned in the literature as EX8, part of an ‘external survey’ (Federici et al. 1990). It was designated EXT8 by Battistini et al. (1993) and lies at a projected distance of 27 kpc from the centre of M 31. Based on an analysis of the spatial distribution and kinematics of GCs in the M 31 halo, EXT8 does not appear to be associated with known M 31 halo substructure (Mackey et al. 2019a). While metallicity variations have been reported due to [Fe/H] = −2.07 for EXT8, previously for EX8 from spectroscopic analyses (Fan et al. 2011; Chen et al. 2016), the HIRES spectrum showed very clearly that the metallicity values are much weaker than in the metal-poor Galactic GC M 15. Most of the α-elements (Si, Ca, Ti) were found to be enhanced by about a factor of two compared to scaled-solar composition, as is typical for metal-poor old stellar populations, but Mg was found to be strongly deficient with [Mg/Fe] = −0.35 ± 0.05 (Larsen et al. 2020). Although the abundances of the α-elements do not vary strictly in lockstep everywhere (Villaume et al. 2020), it is very unusual for Mg to be as discrepant as observed for EXT8, and a similar combination of low [Mg/Fe] and enhancement of other α-elements is not observed in metal-poor Galactic field stars (Frebel & Norris 2015; Kobayashi et al. 2020). One possibility is that the Mg deficiency is related to an extreme case of the Mg−Al anticorrelation and the ‘multiple stellar populations’ phenomenon in GCs (Bastian & Lardo 2018; Gratton et al. 2019).

EXT8 provides us with a unique opportunity to explore the CMD of a stellar population with a metallicity near [Fe/H] = −3 and compare it with stellar model predictions at a metallicity that is about 0.5 dex lower than the current limit set by the metallicity floor of Galactic GCs. The expectation is that it may be an extreme case of the Mg−Al anticorrelation and the ‘multiple stellar populations’ phenomenon in GCs (Bastian & Lardo 2018; Gratton et al. 2019).

2. Observations

EXT8 was observed with the Wide Field Camera 3 (WFC3) on board HST on 13−15 Jan. 2021 (programme ID 16459, PI S. S. Larsen). Exposures were obtained in the filters F300X, F606W, and F814W, using two orbits each for the F300X and F606W observations and three orbits for F814W. The observations were split into two sub-exposures per orbit for a total of four individual exposures in F300X and F606W and six in F814W, all of which were dithered according to the patterns in Anderson (2020). The total exposure times were 5414 s (F300X), 5428 s (F606W), and 8128 s (F814W). EXT8 was placed on the UVIS2 detector and centred near detector coordinates (x, y) = (1040, 1090), with an offset of about 40″ from the centre of the detector towards read-out amplifier C to mitigate the effect of charge-transfer inefficiency. A FLASH = 10 post-flash illumination was further added to the F300X exposures.

The F606W and F814W observations were designed to allow photometry of stars on the upper part of the RGB, approximately down to the level of the HB, while the F300X observations were aimed primarily at detecting and characterising HB and other hot stars. While ultraviolet photometry of RGB stars has been employed extensively to characterise multiple stellar populations in Galactic and extra-galactic GCs (Larsen et al. 2014; Piotto et al. 2015; Niederhofer et al. 2017), reaching the required precision at the distance of M 31 would require much longer exposure times.

Figure 1 shows a colour image of EXT8 based on the F300X, F606W, and F814W exposures, covering a 40″ × 40″ field of view centred on the cluster. It can be seen that EXT8 is well resolved into individual stars, although the cluster is quite compact and crowding becomes severe in the central few arcsec. The cluster is visibly flattened, as was already evident from ground-based imaging (Larsen et al. 2020). Apart from the two bright foreground stars, most stars in the image are members of EXT8 with minimal contamination from the general M 31 halo. Several background galaxies are also visible in the vicinity of EXT8.

3. Analysis

The observations were retrieved from the Mikulski Archive for Space Telescopes (MAST). For our analysis we used the standard pipeline processed images, corrected for charge transfer inefficiencies (the _drc and _flc images). Throughout the paper we assume a distance of 783 kpc for M 31 (Stanek & Garnavich 1998) and a Galactic foreground extinction at the position of EXT8 (J2000.0 coordinates RA = 0°53′14.5″, Dec = +41°33′24.5′′) of A_F300X = 0.353 mag, A_F606W = 0.168 mag, and A_F814W = 0.104 mag in the WFC3 filters (Schlafly & Finkbeiner 2011, via the NASA/IPAC Extragalactic
3.1. Resolved photometry – $F_{606W}$ and $F_{814W}$

We used ALLFRAME (Stetson 1994) to carry out point-spread function (PSF)-fitting photometry of stars in the individual $F_{606W}$ and $F_{814W}$ images (the _flc files). EXT8 is fully contained within the UVIS2 detector and we did not carry out photometry on the UVIS1 detector. The procedure was similar to that adopted in Larsen et al. (2014) and we refer to that paper for more details than those provided here. Briefly explained, the _flc images were multiplied by the pixel area maps available from STScI (Kalirai et al. 2010) and aligned to a common reference frame. As ALLFRAME solves for coordinate transformations between images, the main purpose of the alignment was to enable filtering of bad pixels and cosmic-ray events. Since sub-pixel shifts will tend to blur any sharp features, making them harder to remove, the images were only aligned to within integer pixels shifts. A master-frame was then produced for each filter by average combining the individual frames with the imcombine task in IRAF (Tody 1986), using the ccdclip option to reject pixels that deviated by more than 5σ from the mean at each position. Next, pixels in the individual frames that were flagged as ‘bad’ by imcombine were replaced by their values in the master frame. Stars were detected in the $F_{606W}$ master image with the find task in DAOPHOT, a PSF was produced with the psf task for each individual frame from 36 isolated, bright stars, and a first pass of ALLFRAME photometry was obtained with all $F_{606W}$ and $F_{814W}$ frames as input. Additional stars were detected in a second pass of find on the star subtracted $F_{606W}$ master frame, the PSFs were redetermined for each individual frame on star-subtracted images with only the PSF stars remaining, and a second pass of ALLFRAME photometry was obtained with the combined star list as input. The photometry was calibrated to the standard STMAG system by matching the PSF magnitudes to aperture photometry of the PSF stars in a 10 pixel radius ($0″4$) aperture, applying photometric zero-points computed from the PHOTFLAM header keywords, and adding aperture corrections of $-0.102$ mag ($F_{606W}$) and $-0.107$ mag ($F_{814W}$) from the 10 pixel radius to infinity (Hartig 2009). The star-to-star dispersion of the zero-point offsets between the aperture- and PSF-fitting photometry is about $0.03–0.04$ mag per exposure (Fig. 2), which is consistent with the uncertainties on the individual offsets. The corresponding random uncertainties on the mean zero-point offsets are then less than 0.01 mag per exposure. The photometry from the individual frames was average combined to produce the final photometric catalogue, which is available at the CDS (Table A.1).

The photometric uncertainties and completeness were quantified by means of artificial star experiments. To this end, an artificial cluster was added to each _flc image on the other half of the UVIS2 detector, opposite EXT8. To generate the artificial cluster, artificial stars were sampled at random from a King (1962) profile with a half-light radius of $0″.91$ and a tidal radius of $15″$, which implies a core radius of $0″.21$ (Ishape User’s Guide; Larsen 2014). These core- and half-light radii are similar.
to those found from two-dimensional King profile fits to the images of EXT8 (Sect. 4.4). The tidal radius adopted for these experiments is, however, smaller than the outer radius of EXT8 (about 25″), as a King profile with a larger tidal radius produces too many stars in the outer parts of the cluster. This inconsistency is due to the fact that the actual cluster profile is not well fitted by a King profile in the outer parts (Sect. 4.4). Each artificial star was assigned a mass drawn at random from a Kroupa (2001) mass function, and $F606W$ and $F814W$ magnitudes were then assigned by interpolation in an α-enhanced DSEP isochrone (Dartmouth Stellar Evolution Program; Dotter et al. 2007) with an age of 13 Gyr and [Fe/H] $= -2.5$ (the lowest metallicity available). The total number of artificial stars was scaled such that the integrated magnitude of the artificial cluster was similar to that of EXT8. The DSEP isochrones do not include the HB, but we added a ‘blue sequence’ of artificial stars at $m_{F606W}$ = $m_{F814W}$ = −1.0 and with magnitudes of $m_{F606W}$ = 25.5, 26.0, …, 28.0 to quantify the detection completeness and photometric uncertainties for stars on an extended HB. In each run, five such stars were included per magnitude step, restricting these stars to locations beyond a radius of 3″. The artificial stars were then added to the UVIS2 images with the mksynth task in the BAOLAB package (Larsen 2014) using PSFs obtained with DAOPHOT. A similar number of artificial PSF stars as those used for the photometry of EXT8 were also added to the images. Photometry was then carried out with ALLFRAME in the exact same way as for the analysis of EXT8. To improve statistics, the procedure was repeated 20 times with different random realisations of the artificial cluster and blue sequence, yielding a total of 100 stars for each magnitude step on the blue sequence.

3.2. Resolved photometry – $F300X$

As noted already by Sandage & Walker (1955), the $U$-band brightness of the HB is comparable to that of the RGB. Comparison with BaSTI isochrones (Hidalgo et al. 2018; Pietrinferni et al. 2021) confirms that the brightness of HB stars is expected to remain fairly constant at $M_{F300X} \approx +0.8$ over a range of colours, which is similar to the tip of the RGB in this band. Using the $F300X$ images for source detection is therefore a very efficient way to obtain a fairly clean sample of HB stars, and the entire $F300X$ image in fact contains few stars other than HB stars in EXT8. However, this also introduced some challenges for the photometry by making it difficult to align the individual images and find suitable PSF stars. Moreover, even the HB stars are relatively faint in the $F300X$ images, and requiring a measurement in each individual $F300X$ image would have further reduced the completeness. The photometry involving the $F300X$ images was therefore done separately, following a somewhat different procedure than that outlined above.

For the photometry involving the $F300X$ data we carried out photometry directly on the drizzle-combined images (the _drc files produced by the pipeline). Even so, we could identify only six suitable PSF stars in the $F300X$ image, forcing the assumption of a constant PSF across the UVIS2 detector. It is therefore inevitable that the $F300X$ photometry is of lower quality than that obtained from the $F606W$ and $F814W$ data. The PSFs for the drizzled $F606W$ and $F814W$ images were determined separately using a larger number of PSF stars, as for the reduction based on the _flc images, and all three _drc images were fed to ALLFRAME and analysed simultaneously. The photometry on the drizzled images is available in Table A.2.

A separate set of completeness tests were performed for the photometry involving the $F300X$ image, mostly following the procedure outlined above for the $F606W/F814W$ photometry. As before, a sequence of blue stars with $m_{F606W} - m_{F814W} = -1.0$ and $m_{F606W}$ between 25.5 and 28.0 were added to the images. However, the detection completeness here depends much more strongly on the brightness in the $F300X$ images and to quantify this we therefore also repeated the completeness tests for a range of $F300X$ magnitudes, $m_{F300X} = 25.0, 25.5, …, 27.0$.

3.3. Colour-magnitude diagrams for EXT8 and the artificial cluster

Figure 3 shows the CMDs for $F606W/F814W$ photometry of EXT8 (panel A), a field region centred on the position at which
the artificial cluster was added in the artificial star experiments (panel B), and the combined CMD for the 20 artificial cluster realisations (panel C). In this figure, and throughout the remainder of this paper, photometry is shown for distances in the range $3'' < r < 24''$ from the centre of EXT8 (or the artificial cluster). At radii less than $3''$, the stellar density is too high to obtain reliable photometry and at radii greater than about $25''$ very few cluster members remain. The RGB and HB of EXT8 are clearly visible in panel A and will be discussed in more detail below (Sect. 4).

The solid red curve in the top and centre panels is a ridge line for the upper EXT8 RGB with dashed lines offset by ±0.07 mag in $m_{F606W}$ and $m_{F814W}$. The ridge line is a fourth-order polynomial fit to the colours of the RGB stars brighter than $m_{F606W} = 25.5$. To define the ridge line we used an iterative procedure to reject stars that deviated by more than 0.1 mag from the polynomial fit. Since all panels in Fig. 3 cover the same area on the sky, it can be seen from a direct comparison of panels A and B that contamination from the general M 31 halo field is expected to be very minimal for stars on the upper RGB of EXT8, while the redder part of the HB may be contaminated by a few non-members. A few stars located on the red side of the RGB in panel A also have counterparts in panel B and are likely field stars.

The ridge line in panel C was obtained in the same way as that in panel A, but here from a fit to the artificial star photometry. It coincides closely with the isochrone used to generate the input photometry for the artificial star experiments (white dashed curve). The coloured circles indicate the detection completeness determined for the sequence at $m_{F606W} = m_{F814W} = -1.0$ and for stars along the isochrone in 0.5 mag bins of $m_{F606W}$. The 50%, 90%, and 97% completeness levels determined for stars on the blue sequence and the isochrone are connected with dashed lines. For both the blue sequence and the isochrone, the completeness at the faint limit of the plot range, $m_{F606W} = 28$, is about 38%. The 90% completeness limit is at $m_{F606W} \approx 26.8$ and for RGB stars brighter than $m_{F606W} = 25.5$ the completeness is greater than 97%. As these completeness fractions were determined from the artificial cluster, they represent an overall level of completeness for all stars with $r > 3''$.

In addition to the artificial stars, the real stars from panel B are again visible in panel C. Due to small variations in the PSF reconstruction and the photometric zero-point calibrations in the 20 realisations of the artificial star photometry, the magnitudes and colours of these stars differ slightly between the 20 realisations. The dispersions of the magnitudes and colours are less than 0.005 mag, which provides an additional estimate of the internal precision of the photometric calibration. However, the total uncertainty is likely somewhat larger.

### 3.4. Photometric uncertainties

In Fig. 4 we compare the input colours and magnitudes for the artificial cluster stars with the ALLFRAME measurements. The small grey points show the differences between measured and input values, $\Delta_{\text{ALLFRAME-In}}$, for stars sampled from the isochrone, and larger light blue dots are for stars on the blue sequence. The solid black and blue lines show the median differences in bins of 0.5 mag with error bars determined from 100 Monte Carlo experiments. For stars on the isochrone, the error bars on the median values are generally smaller than the width of the black line. The sharp feature at $m_{F606W} = 23$ is due to the artificial PSF stars while the fainter feature near $m_{F606W} \approx 24.2$ is a real feature of the RGB luminosity function (LF), the RGB bump (Bjork & Chaboyer 2006). The distribution of $\Delta_{\text{ALLFRAME-In}}$ versus input magnitude is generally very minor. For stars brighter than the HB level, $m_{F606W} \approx 25$, the median magnitude offsets are between $-0.003 \pm 0.001$ mag and $-0.007 \pm 0.002$ mag (top panel) and the colour offsets are $\approx (0.007 \pm 0.001)$ mag (bottom panel). The zero-point of the input magnitudes are somewhat poorly defined because the empirical PSFs used to generate the artificial stars are imperfect approximations to the true PSF. This likely accounts for any small overall differences in the zero-points of the input- and measured magnitudes. The median offsets remain small for stars sampled from the isochrones also at fainter magnitudes, reaching $-0.043 \pm 0.004$ mag in $\Delta_{\text{ALLFRAME-In}}$ and $-0.004 \pm 0.003$ mag in $\Delta_{\text{ALLFRAME-In}}$ at $m_{F606W} = 27.5$. For the blue sequence the uncertainties are larger because of the smaller numbers of artificial stars, but the median $\Delta_{\text{ALLFRAME-In}}$ values are mostly consistent with those seen for the stars sampled from the isochrones.

In Fig. 5 we compare three different ways of estimating the photometric uncertainties: $\sigma_{\text{ALLFRAME}}$ was obtained by propagating the errors on the individual measurements reported by ALLFRAME, $\sigma_{\text{CMD}}$ was computed from the frame-to-frame dispersions of the magnitude measurements for each star, and $\sigma_{\text{Yon}}$ is the dispersion of the $\Delta_{\text{ALLFRAME-In}}$ values. In general, the $\sigma_{\text{ALLFRAME}}$ and $\sigma_{\text{CMD}}$ estimates are very similar, which shows that the errors reported by ALLFRAME accurately represent the random uncertainties on the photometry. However, both are significantly

![Fig. 4. Difference between input and ALLFRAME measurements of $m_{F606W}$ magnitudes (top) and $m_{F606W}-m_{F814W}$ colours (bottom) for artificial star experiments. Solid black lines show median values for artificial stars sampled from isochrones (grey points) while dashed blue lines show median values for artificial stars on the blue sequence (blue dots).](image-url)
Fig. 5. Comparison of error estimates for artificial star tests: $\sigma_{\text{syn}}$ are the dispersions of the ALLFRAME-input magnitudes, $\sigma_{\text{ALLFRAME}}$ are the ALLFRAME errors propagated to the mean magnitudes, and $\sigma_{\text{ff}}$ are the errors on the mean ALLFRAME magnitudes as determined from the frame-to-frame dispersion. Solid lines are for stars sampled from the isochrone while dashed lines are for stars on the blue sequence.

smaller than $\sigma_{\text{syn}}$. The difference is most pronounced for the $m_{F606W}$ magnitude measurements, where $\sigma_{\text{syn}}$ is about three times greater than $\sigma_{\text{ALLFRAME}}$ and $\sigma_{\text{ff}}$ (top panel). For the $m_{F606W} - m_{F814W}$ colours the agreement is better, but the $\sigma_{\text{syn}}$ estimates remain somewhat larger than the ALLFRAME errors also here. These differences between the errors based on the photometry alone and those based on the comparison with artificial star tests are most likely due to the effect of unresolved blends and crowding, which can introduce a bias on the magnitude measurements without necessarily affecting the random uncertainties strongly. Indeed, we found that the $\sigma_{\text{syn}}$ values agreed more closely with $\sigma_{\text{ALLFRAME}}$ and $\sigma_{\text{ff}}$ for stars further from the centre of the artificial cluster. For example, at radii $> 8''$ the $\sigma_{\text{syn}}$ errors on the $m_{F606W}$ magnitudes exceed $\sigma_{\text{ALLFRAME}}$ by only $\approx 50\%$. We also note that all three uncertainty estimates agree well for the bin at $m_{F606W} = 23$. This bin contains the artificial PSF stars, which were deliberately placed in uncrowded parts of the image. We further note that the uncertainties on the $m_{F606W}$ magnitudes are similar for stars sampled from the isochrone and from the blue sequence, whereas the uncertainties on the colours are greater for stars on the blue sequence. The latter is explained by the fact that stars on the blue sequence are fainter in the $F814W$ images.

In summary, the ALLFRAME photometry reproduces the input colours of the artificial stars without significant biases ($< 0.01$ mag) across the relevant parts of the CMD. The full uncertainties are, however, larger than the formal errors computed by ALLFRAME and are therefore better quantified using the artificial star experiments.

4. Results

In Fig. 6 we show the CMD of EXT8 together with those of the Galactic GCs M 15, M 92, M 30, NGC 4147, and M 5 (in order of increasing metallicity). These clusters have been selected to have relatively small foreground extinctions, so as to reduce uncertainties in the comparison of the CMDs. Their distances, foreground extinctions, and metallicities are listed in Table 1. M 15, M 30, and M 92 are among the most metal-poor GCs in the Milky Way, while NGC 4147 and M 5 are about 0.5 dex and 1.0 dex more metal-rich, respectively. The composition of NGC 4147 has been less extensively studied than those of the other clusters, but a recent determination of its metallicity ([Fe/H] = $-1.84 \pm 0.02$; Villanova et al. 2016) agrees well with the value in the Harris (1996) catalogue given in the table. The photometry for the Galactic GCs was transformed from the catalogues of $F606W$ and $F814W$ magnitudes provided by the ACS Galactic GCs Survey (ACSGCS; Sarajedini et al. 2007; Anderson et al. 2008) to the WFC3 experiments using the relations in Deustua & Mack (2018). The transformed VEGAMAG magnitudes from the ACSGCS data were further converted to the STMAG system by adding the differences between the corresponding zero-points ($\Delta_{F606W} = 0.246$ mag and $\Delta_{F814W} = 1.259$ mag), taken from the WFC3/UVIS web pages at STScI.

For each Galactic GC we show an RGB ridge line together with the ridge line for EXT8, again obtained as polynomial fits. We discuss the RGBs in more detail below (Sect. 4.1) but it is already clear from Fig. 6 that the differences below [Fe/H] = $-2$ are fairly minor. The EXT8 RGB is very similar in colour to that of M 15, and both are slightly bluer than the M 92 and M 30 RGBs. For NGC 4147 and, especially, for M 5, the RGBs are noticeably redder, as expected for their higher metallicities. Visually, the EXT8 ridge line appears slightly steeper than for the other GCs.

The HB will be discussed in Sect. 4.2. To first order, the HB morphology of EXT8 resembles those of M 30, M 92, and NGC 4147. The detected HB stars lie mostly on the blue side of the instability strip (IS), the edges of which are marked on the EXT8 CMD in Fig. 6. A few stars are located within the IS; these will be considered further in Sect. 4.2.3. There is no evident excess of HB stars on the red side of the IS, although it cannot be excluded that a few red HB stars are present. The edges of the IS were calculated using the relations in Marconi et al. (2015), with some extrapolation to the lower metallicity. The blue and red edges are at $T_{\text{eff}} \approx 7475$ K and $T_{\text{eff}} \approx 5888$ K, which corresponds to $M_{F606W} - M_{F814W} = -0.8$ and $M_{F606W} - M_{F814W} = -0.5$ from interpolation in the colour–$T_{\text{eff}}$ relation for stars on the HB of an [Fe/H] = $-2$ BarSTI isochrone. It is not obvious from Fig. 6 whether the HB of EXT8 has an extended blue part, as in M 15, and we will address this point in more detail based on the $F300X$ data in Sect. 4.2.2.

The dynamically determined mass-to-light ratio and the hydrogen Balmer lines already strongly imply that EXT8 is an old, metal-poor GC (Fan et al. 2011; Chen et al. 2016; Larsen et al. 2020), and this is further confirmed by the general morphology of its CMD.

4.1. The red giant branch

4.1.1. Comparison with Milky Way GCs

To quantify the differences between the RGB colours of the various clusters, we first measured the mean colour offsets with respect to the M 92 RGB ridge line, $\Delta(M_{F606W} - M_{F814W})_{M 92}$, in two bins of absolute magnitude (Table 2). The brighter bin

\[ \Delta(M_{F606W} - M_{F814W})_{M 92} \]
One might wonder how robust the RGB colour offsets with respect to M 92 ridge line. The RGBs of both clusters are about 0.03 mag bluer than those of M 30 and M 92. From the same ACSGCS photometry used here, Vandenberg et al. (2013) found that M 15 also has slightly bluer main sequence turn-off (MSTO) colours compared to M 30 and M 92. They pointed out that a decrease in the reddening correction towards M 15 of about 0.02 mag in $E(B-V)$ (also corresponding to a difference in $E(F606W-F814W) \approx 0.02$) would bring the MSTO colours of the three clusters into closer agreement. Given the higher foreground reddening towards M 15, it may be reasonable to suspect that the uncertainty on its reddening correction is somewhat larger. A small redwards shift in the colours of M 15 would then make EXT8 bluer than all three metal-poor Galactic GCs.

One might wonder how robust the RGB colour offsets measured for EXT8 with respect to the Galactic GCs are. (−2.5 < $M_{F606W}$ < −1.0) mostly lies above luminosities where the asymptotic giant branch (AGB) may bias the measurements of RGB colours (see for example the CMDs for M 15 and M 92), while the fainter bin (+1.0 < $M_{F606W}$ < +1.5) lies below the HB. EXT8 does indeed have a bluer RGB than most of the Milky Way GCs, although M 15 has colour offsets nearly identical to those of EXT8. The RGBs of both clusters are about 0.03 mag bluer than those of M 30 and M 92. We note that the brighter bin has a small offset even for M 92 itself, which illustrates the uncertainties related to the different ways of quantifying the RGB colours (polynomial fits versus average colours of stars in some magnitude range).

Given the low metallicity of EXT8, it is not surprising that it should have a very blue RGB. The reason why the RGB of M 15 matches that of EXT8 more closely than those of M 30 and M 92 is less clear. Given the similar metallicities of M 15, M 30, and M 92. From the same ACSGCS photometry used here, Vandenberg et al. (2013) found that M 15 also has slightly bluer main sequence turn-off (MSTO) colours compared to M 30 and M 92. They pointed out that a decrease in the reddening correction towards M 15 of about 0.02 mag in $E(B-V)$ (also corresponding to a difference in $E(F606W-F814W) \approx 0.02$) would bring the MSTO colours of the three clusters into closer agreement. Given the higher foreground reddening towards M 15, it may be reasonable to suspect that the uncertainty on its reddening correction is somewhat larger. A small redwards shift in the colours of M 15 would then make EXT8 bluer than all three metal-poor Galactic GCs.

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Anderson et al. (2008) found that ’unmodellable’ PSF variations due to focus changes, geometric distortions, and detector inhomogeneities can introduce zero-point uncertainties of up to 0.02 mag across the field-of-view for the ACSGCSC photometry of Galactic GCs. The WFC3 is not immune to such effects (Sabbi & Bellini 2013), and while Anderson et al. (2008) had access to large numbers of bright PSF stars across the field of view, we are forced to rely on fainter stars, so that our PSFs are necessarily of lower fidelity. While the scatter in the photometric zero-points among the calibration stars is largely consistent with the individual 0.03–0.04 mag uncertainties (Sect. 3.4), this could hide systematic effects at the level of ~0.01 mag (Fig. 2).

It therefore cannot be excluded that zero-point uncertainties of 0.01–0.02 mag are present in our photometry of EXT8. If the reddening of EXT8 is lower than we have assumed, this could in principle also produce a colour shift. However, a shift of 0.03–0.04 mag relative to the value of $E(F606W−F814W) = 0.064$ adopted here would require the reddening to be overestimated by about a factor of two, which seems unlikely. Indeed, the Schlafly & Finkbeiner (2011) reddening at the position of EXT8 is $E(B−V) = 0.06$ mag, which is similar to or slightly less than the reddenings found for other GCs around M 31 from CMD analyses (Mackey et al. 2006, 2007). Differences between the ACS and WFC3 systems are probably only a minor contributor to the uncertainties. The ACS and WFC3 F814W systems are essentially identical (Deustua & Mack 2018), while the conversion between the ACS and WFC3 F606W systems involves a slight colour dependence of up to 0.02 mag for a VEGAMAG colour of $m_{F606W,Vega} = m_{F814W,Vega} = 1$, typical of RGB stars. After accounting for this, as we also do here, Deustua & Mack (2018) found consistency between the photometric calibrations of ACS and WFC3 at the 0.5% level, corresponding to an uncertainty of 0.005 mag. We also recall (Fig. 4) that systematic biases in the ALLFRAME photometry as a function of magnitude are below the 0.01 mag level for RGB stars. In summary, then, it seems unlikely that the colour difference between the EXT8 RGB and the metal-poor Galactic GCs can be fully explained by uncertainties in the measurements, so that a real difference is likely present.

Differences in metallicity are expected to be accompanied not only by a shift in mean RGB colours but also by a change in the slope of the RGB (Sandage & Wallerstein 1960; Demarque & Geisler 1963; Hartwick 1968; Saviane et al. 2000; Streich et al. 2014). From a practical point of view, the change in the slope of the RGB (Sandage & Wallerstein 1960; Streich et al. 2014), and BaSTI isochrones. For clarity the other MIST isochrones are only available for scaled-solar composition. They are therefore shown for the [Fe/H] = −2.5 but they are otherwise shown for the same [Fe/H] values as the BaSTI isochrones. Unlike the DSEP and BaSTI isochrones, which are both shown for a-enhanced composition ([α/Fe] = +0.4), the MIST isochrones are only available for scaled-solar composition. They are therefore shown for [Fe/H] = −3.0, −2.3, −2.0, −1.7, and −1.1, which correspond to about the same total metallicities as those for which the BaSTI and DSEP isochrones are plotted (see, for example, Salaris et al. 1993). All isochrones are shown for an age of 13 Gyr. Apart from the fact that no set of isochrones perfectly matches all of the observed ridge lines, perhaps the most obvious difference is the much wider colour separation between the MIST isochrones at low metallicities compared to the BaSTI set. The [Fe/H] = −3.0 and −2.3 MIST isochrones are separated by about 0.08 mag in $M_{F606W−F814W}$ near the tip of the RGB, while the two corresponding BaSTI isochrones ([Fe/H] = −3.2 and −2.5) are indistinguishable in the figure. Overall, the BaSTI isochrones provide the best fits to the ridge lines for the range of magnitudes and metallicities shown here. The DSEP isochrones

\[ M_{F606W−F814W} = −0.168 \text{ and } −0.313, \text{ and the resulting RGB slope is } \Delta(M_{F606W−F814W})/\Delta(M_{F606W−M_{F814W}}) = −13.8±0.9. \]
tend to be somewhat too red (as noted by Vandenberg et al. 2013), and the slopes of the most metal-poor MIST isochrones are shallower than the observed ridge lines. The difference in RGB slope between metal-poor MIST and BaSTI isochrones was previously noted by Hidalgo et al. (2018).

A different way of comparing the various isochrones is shown in Fig. 9, where the colours at a magnitude of $M_{F606W} = -1.5$ are plotted as a function of metallicity for the different models. We also plot colours for 10 Gyr old BaSTI isochrones, which are about 0.01 mag bluer at a given metallicity compared to the 13 Gyr models. The effect of age differences of a few Gyr is thus relatively small, albeit not entirely negligible, compared to the small RGB colour differences discussed here. We include the colours of the RGB ridge lines for all GCs from Fig. 6 and we have drawn M 15 with an open symbol as a reminder that its reddening correction may be more uncertain than for the other clusters. The larger error bar for NGC 4147 is due to its peculiarly low magnesium abundance of EXT8. Vandenberg et al. (2012) investigated the effect of varying the abundances of specific elements and found that an increase of 0.4 dex in [Mg/Fe] made the upper RGB of an [Fe/H] = −2 isochrone (at $M_V = −2$) about 0.05 mag redder in $V − I$ (read off of their Fig. 16). We compared the $V − I$ colours of 13 Gyr, [Fe/H] = −1.9 BaSTI isochrones with scaled-solar composition and [α/Fe] = +0.4 and found the difference to be about 0.04 mag, which is comparable to the effect of varying Mg alone. While a change in Mg abundance can clearly have a non-negligible effect on the colours, the effect is likely smaller at the lower metallicity of EXT8. Given that the effect of changing the Mg abundance appears to be comparable to that of changing [α/Fe] as a whole, an estimate may be made by comparing the colours of scaled-solar and α-enhanced BaSTI isochrones. For [Fe/H] = −2.9 and an age of 13 Gyr we found a colour difference in $M_{F606W} − M_{F814W}$ of only 0.005 mag at $M_{F606W} = −1.5$. Variations in helium abundance, as may be associated with multiple populations (Milone et al. 2018), will also affect the RGB
do show a continued trend towards bluer colours, but the predicted effect now appears too large compared with the observations. The DSEP isochrones are again redder than the MIST and BaSTI isochrones as well as the observed RGBs.

An extensive discussion of the physical differences between the various isochrones would lead too far here. We note, for example, that the BaSTI isochrones are available for different assumptions about overshooting, mass loss, and diffusion, which can lead to differences of about 0.01 mag in the RGB colours at [Fe/H] = −2.9. A more detailed exploration of how these and other differences in the underlying physics affect the behaviour of the models in the very metal-poor regime may be worthwhile. In addition to differences in the physical parameters of the isochrones, significant differences in the observational plane can also arise depending on whether synthetic or empirical colour transformations are used (Dotter et al. 2007).

When discussing RGB colour differences at the level of a few times 0.01 mag, higher-order differences in the chemical abundance patterns may become relevant. In particular, we recall the peculiarly low magnesium abundance of EXT8.
colours. According to the BaSTI isochrones, an increase from $Y = 0.247$ to $Y = 0.30$ at [Fe/H] = −2.9 makes the RGB about 0.004 mag bluer, so the effect is again expected to be small. While the different behaviour of the various isochrones at these low metallicities must be kept in mind, it appears unlikely that the detailed abundance patterns of EXT8 have a strong effect on the comparisons in the preceding paragraphs.

4.1.3. Constraints on the metallicity spread

To constrain any metallicity spread in EXT8, we compared the observed colour distribution of its RGB stars with the artificial star experiments. For this comparison we included stars down to a magnitude limit of $M_{F606W} = 24.5$, located within the dashed lines in Fig. 3. At fainter magnitudes, the isochrones lie closer together and the observational errors increase, so that the sensitivity to a metallicity spread diminishes. Figure 10 shows the distributions of offsets $\Delta_{F606W-F814W}$ between the measured colours of RGB stars and the ridge line for both the observed CMD and the artificial star experiments. For the observed CMD the dispersion is $\sigma_{\text{obs}} = 0.0197$ mag (for 79 stars), and for the artificial stars we find $\sigma_{\text{syn}} = 0.0185$ mag (1773 stars). The observed dispersion is thus slightly larger than that indicated by the artificial star experiments. Assuming that the $\Delta_{F606W-F814W}$ offsets follow a normal distribution we can calculate the probability of finding a dispersion equal to or greater than $\sigma_{\text{obs}} = 0.0197$ mag when drawing 79 stars from a parent distribution with a dispersion of $\sigma_{\text{syn}} = 0.0185$ mag from the $x^2$ statistics. This probability is $P(\sigma_{\text{syn}} > \sigma_{\text{obs}}) = 23\%$, so there is no strong evidence of the difference being significant, even if we assume that the artificial star tests perfectly capture the observational uncertainties.

As an additional test, we carried out a series of Monte Carlo tests in which we resampled 79 $\Delta_{F606W-F814W}$ offsets multiple times from the artificial star photometry and recomputed the dispersion. The dispersions exceeded the observed value of 0.0197 mag in about 21% of the realisations, in good agreement with the $x^2$-based estimate. The artificial star CMD, which is based on DSEP isochrones, does not include any potential contamination of the RGB by AGB stars. We repeated the artificial star experiments using BaSTI isochrones, which do include the AGB, and found that the dispersion of the $\Delta_{F606W-F814W}$ distribution indeed increased slightly to $\sigma_{\text{syn}} = 0.0219$ mag, now slightly broader than $\sigma_{\text{obs}}$. We are led to conclude that the observed colour spread on the RGB can be fully explained by the photometric uncertainties.

We next quantify how large a metallicity spread can be ruled out. This is complicated by the significantly different colour–metallicity relations predicted for the RGB by the various isochrones, and we therefore started with a more model-independent comparison based on colour spreads alone. To this end we randomly added normally distributed colour offsets to the $\Delta_{F606W-F814W}$ offsets from the artificial star tests. We then compared the resulting colour distributions with the observations to find out whether the broadened artificial star colour distribution was still consistent with being at least as narrow as the observed colour distribution. For an intrinsic colour spread of $\sigma_i = 0.005$ mag, the total dispersion of the artificial star colour distribution was 0.0193 mag, still slightly narrower than the observed distribution, and the probability of finding a dispersion smaller than or equal to the observed dispersion was found to be $P(\sigma_{\text{syn}} < \sigma_{\text{obs}}) = 67\%$, implying that this is certainly allowed. For larger intrinsic colour dispersions, the corresponding probabilities were found to be $P(\sigma_{\text{syn}} < \sigma_{\text{obs}}) = 27\%$ (for $\sigma_i = 0.010$ mag), $P(\sigma_{\text{syn}} < \sigma_{\text{obs}}) = 2.4\%$ ($\sigma_i = 0.015$ mag), and $P(\sigma_{\text{syn}} < \sigma_{\text{obs}}) = 0.04\%$ ($\sigma_i = 0.020$ mag). An intrinsic colour dispersion greater than 0.015 mag is therefore ruled out at about the two sigma level. When basing the analysis on the BaSTI artificial star experiments, this limit is reduced to about 0.010 mag.

To map these colour dispersions to metallicity dispersions, we used the isochrones for guidance. Interpolating between the BaSTI [Fe/H] = −2.5 and [Fe/H] = −3.2 isochrones, we found an RGB colour of $M_{F606W} - M_{F814W} = -0.184$ for [Fe/H] = −2.9 at a reference magnitude of $M_{F606W} = -1.5$ (Fig. 9). The [Fe/H] = −2.2 isochrone has $M_{F606W} - M_{F814W} = -0.169$, about 0.015 mag bluer. This is almost the same as the 2$\sigma$ limit on the colour dispersion, so that a metallicity dispersion larger than the difference between the two isochrones, or 0.7 dex, is then ruled out at about the 2$\sigma$ confidence level. However, this assumes that a symmetric metallicity distribution maps to a symmetric colour distribution, which is probably not the case according to the BaSTI isochrones – indeed, the lower bound on [Fe/H] is essentially unconstrained since there are no BaSTI isochrones with [Fe/H] < −3.2, but we recall that the RGB colours remain nearly constant between [Fe/H] = −2.5 and [Fe/H] = −3.2.

If we instead use the MIST isochrones, interpolation gives $M_{F606W} - M_{F814W} = -0.219$ for [Fe/H] = −2.7 (accounting for the scaled-solar composition). Owing to the wider colour separation between these isochrones, a metallicity dispersion of about 0.2 dex maps to a colour dispersion of about 0.013 mag in both directions, again not far from the 0.015 mag upper limit. Were we to rely on the MIST isochrones, the two-sigma upper limit on the metallicity dispersion would therefore be about 0.2 dex.

As an alternative to these model comparisons, we may take the $M_{F606W} - M_{F814W}$ colour difference of −0.03 mag between the RGB of EXT8 and the metal-poor Galactic GCs (Sect. 4.1.1) as indicative of a metallicity difference of −0.6 dex. The allowed −0.015 mag colour dispersion would then translate to a metallicity dispersion of $\sigma_{[\text{Fe/H}]} \approx 0.3$ dex. In summary, the main uncertainty in quantifying the limit on the metallicity spread is the colour–metallicity relation for the metal-poor RGB stars.
4.1.4. The tip of the RGB and the distance to EXT8

The $I$-band luminosity of the tip of the RGB (TRGB) is only weakly sensitive to metallicity and is commonly used as a distance indicator (Lee et al. 1993). Although the TRGB is not well sampled in our CMD of EXT8, we can use the brightest RGB star (marked with a circle in Fig. 1) to obtain a consistency check of the adopted distance, assuming this star is indeed a member of EXT8. The star is located well outside the crowded central region of EXT8 and has reddening-corrected magnitudes of $m_{F606W,0} = 21.71 \pm 0.01$ and $m_{F814W,0} = 21.66 \pm 0.01$. Transforming these magnitudes to the Johnson–Cousins system (Harris 2018) gives $m_{I,0} = 20.40$ and $(V - I)_0 = 1.32$, or $M_I = -4.07$ for the assumed distance. Recent calibrations of the TRGB magnitude range between $M_{I,TRGB} = -3.99$ and $M_{I,TRGB} = -4.16$ (Jang & Lee 2017; Capozzi & Raffelt 2020). While these calibrations are given for somewhat redder TRGB colours, the corresponding $M_{I,TRGB}$ magnitudes would be less than 0.01 mag fainter at the colour of the EXT8 TRGB (Jang & Lee 2017). The photometry of the brightest RGB star is thus consistent with EXT8 being located at the assumed distance.

However, given that the upper part of the RGB is not well sampled, it is quite possible that we have not detected the true tip of the RGB, in which case the agreement would be less good and EXT8 might be somewhat closer than assumed. Indeed, our CMD only contains 14 stars within one magnitude of the brightest RGB star, while a reliable estimate of the TRGB magnitude requires at least ~50 stars in this range (Madore & Freedman 1995). The TRGB then appears consistent with a slightly smaller distance, as also suggested from isochrone fitting to the HB, but does not favour the larger distance modulus suggested by the comparison of the HB brightness with the empirical relation of Dotter et al. (2010; Sect. 4.2). If EXT8 is indeed closer than assumed, this would make the RGB redder at a given absolute magnitude. If we decrease the distance modulus by 0.1 mag and repeat the comparison of colour offsets from Sect. 4.1.1, we find a shift of 0.01 mag towards the red at fixed $M_{F606W}$. The RGB of EXT8 thus remains bluer than all of the Galactic GCs except M 15 for any reasonable range of distances allowed by the data. It would take a shift of about 0.3 mag in the distance modulus (100 kpc in the distance) to make the EXT8 RGB as red as those of M 92 and NGC 5466, but such a large shift would make the HB clearly too faint.

4.2. The horizontal branch

4.2.1. General morphology and HB magnitude

Classical parameters to quantify HB morphology include the Mironov (1972) index, $M_I \equiv B/(B + R)$, and HBR $= (B - R)/(B + V + R)$ (Lee et al. 1994), where $B$, $V$, and $R$ are the numbers of stars bluer than the IS, in the IS, and redder than the IS, respectively. The apparent simplicity of these definitions is somewhat deceptive and complications arise because (1) not all HB stars may be detected, in particular if a blue tail is present and extends below the detection limit, (2) RR Lyrae stars may scatter out of the IS if the pulsation periods are incompletely sampled, and (3) red HB stars may be confused with RGB stars. For GCs at the distance of M 31, where RR Lyrae identification is challenging, a simplified version of the Mironov index (SMI) can be defined by splitting the HB in two, roughly in the middle of the IS (Rich et al. 2005). Here we adopt the definition of the SMI from Perina et al. (2012), counting red HB (RHB) stars as stars having $0.5 < (V - I)_0 < 0.8$ and lying within $0.5 \text{ mag}$ of the $V$ magnitude of the HB ($V_{\text{HB}}$) and blue HB (BHB) stars as stars having $(V - I)_0 < 0.5$ and with $V_{\text{HB}} - 1 < V < 26$. Converted to WFC3 STMag magnitudes, this corresponds to $-0.59 < M_{F606W} - M_{F814W} < -0.33$ for the colour limits of the RHB and $M_{F606W,0} - 1 < m_{F606W,0} < 26.25$ for the magnitude limits of the BHB.

To determine the magnitude of the HB, $m_{F606W,HB,0}$, we followed a procedure similar to that described in Dotter et al. (2010). Noting that the HB of M 5 spans a wide range of colours, they shifted the HBs of other GCs relative to the M 5 HB and then determined $m_{F606W,HB}$ as the $m_{F606W}$ magnitude of the horizontal part of the shifted M 5 HB. For M 5 itself they found $m_{F606W,HB,VEGA} = 14.90 \pm 0.05$ in the VEGAMAG system. We defined a fiducial HB sequence for M 5 using the same ACS/GCS data as Dotter et al. (2010; Fig. 6). We calculated the median $m_{F606W}$ magnitude in 0.05 mag bins of $M_{F606W} - M_{F814W}$ colour for $-1.10 < M_{F606W} - M_{F814W} < -0.45$, skipping the range $-0.8 < M_{F606W} - M_{F814W} < -0.5$ where RR Lyrae variables cause a large scatter in colour and magnitude in the M 5 CMD. For the resulting fiducial sequence we found $m_{F606W,HB,0} = 15.07$ at $M_{F606W} - M_{F814W} = -0.5$. Converting to VEGAMAG magnitudes and adding back the foreground extinction, this becomes $m_{F606W,HB,VEGA} = 14.92$, which matches the value found by Dotter et al. (2010) within 0.02 mag.

Figure 11 shows the reddening corrected CMD of EXT8 with the shifted M 5 HB sequence and the BHB and RHB boxes overlapped. A shift of $\Delta m_{F606W} = 9.88 \pm 0.03$ mag was required to match the fiducial M 5 HB to the EXT8 HB, calculated for colours in the range $-1.0 < M_{F606W} - M_{F814W} < -0.8$. This gave $m_{F606W,HB,0} = 24.96 \pm 0.03$ mag. While, as noted above, a few RHB stars may be present in the CMD, it is obvious that the vast majority of the 55 stars located within the RHB box are RGB stars. In a few similar cases, Perina et al. (2012) were able to detect a peak in the RGB LF near the expected level of the RHB and they could then statistically subtract the underlying RGB population via interpolation in the LF. However, no such excess of RHB candidates was detected for EXT8. The BHB

Fig. 11. Colour–magnitude diagram of EXT8 overlaid with the HB of M 5 and an [Fe/H] = -2.9 BaSTI isochrone with an age of 13 Gyr (in both cases shifted vertically to match up with EXT8). The boxes show the selection regions for BHB and RHB stars, and the slanted lines show the IS boundaries. The two RR Lyrae candidates are marked with larger red circles.

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The assumed distance, about 0.1 mag fainter than the expected relation in Dotter et al. gives $M_{\text{M5}} = 2005$; Federici et al. 2012). Extrapolated to $[\text{Fe}/\text{H}] = -2.9$, the relation in Dotter et al. gives $M_{\text{M606}} = +0.14 \pm 0.04$ or $M_{\text{M606}} = +0.39 \pm 0.04$ (STMag). The value measured for EXT8 from the shifted M 5 HB is $+0.49 \pm 0.03$ at the assumed distance, about 0.1 mag fainter than the expected value. However, at $[\text{Fe}/\text{H}] = -2.4$ the predicted magnitude is $M_{\text{M606}} = +0.50$, which is nearly identical to the value measured for EXT8. If EXT8 is indeed located at the assumed distance, this might suggest that the $M_{\text{M606}}$ versus $[\text{Fe}/\text{H}]$ relation flattens below $[\text{Fe}/\text{H}] = -2.4$. Alternatively, taking the predicted $M_{\text{M606}}$ for $[\text{Fe}/\text{H}] = -2.9$ at face value, the implied distance modulus would be $(m - M)_0 = 24.57 \pm 0.05$, taking into account the uncertainties on both the predicted and measured HB magnitudes.

In the same way that the M 5 HB fiducial sequence was shifted to match the EXT8 HB, we determined the vertical shift that gave the best fit to a theoretical isochrone. Figure 11 includes a 13 Gyr BaSTI isochrone interpolated to $[\text{Fe}/\text{H}] = -2.9$, for which we found a shift of $(m - M)_0 = 24.43 \pm 0.03$ mag. This is consistent with the distance modulus $(m - M)_0 = 24.47$ that corresponds to the assumed distance of M 31, but we note that this small shift is in the opposite sense to that obtained from the Dotter et al. (2010) relation and would thus lead to a larger discrepancy between the measured HB magnitude and the prediction. The independent constraint on the distance obtained from the tip of the RGB also appears to support a distance to EXT8 no greater than the assumed mean distance of M 31 (Sect. 4.1.4). The shape of the HB predicted by the BaSTI models is somewhat different from the M 5 fiducial sequence, probably due to the higher metallicity of M 5. Pietrinferni et al. (2021) show that the HBs of the BaSTI isochrones match GCs of different metallicities well, and also reproduce the trend of HB luminosity versus metallicity. Part of the discrepancy between the predicted and observed $M_{\text{F606}}$ magnitudes might then be the result of applying the M 5.5 HB at the low metallicity of EXT8.

Our estimates of $(m - M)_0$ thus lie in the range between 24.43 $\pm$ 0.03 (about 770 kpc) and 24.57 $\pm$ 0.05 (824 kpc), with a preference for distances towards the lower end of the range. This is then consistent with a distance modulus of about 24.46 for M 31 (Stanek & Garnavich 1998; Conn et al. 2012) but still allows a considerable range of line-of-sight separations between the centre of M 31 and EXT8. It is also worth noting that the distance of M 31 itself is subject to systematic uncertainties, depending on the choice of distance indicators, where a shorter distance modulus of $(m - M)_0 = 24.38 \pm 0.06$ (752 kpc) has been deduced from Cepheids (Riess et al. 2012). The range of three-dimensional separations between EXT8 and the centre of M 31 could then lie between the projected value of 27 kpc and up to more than 70 kpc, although the latter would require a perhaps somewhat unlikely conspiracy of uncertainties.

4.2.2. The HB in the ultraviolet

Based on the artificial cluster CMD and completeness tests (Fig. 3), HB stars in EXT8 are expected to be detectable well below the boundaries of the BHB box. However, at magnitudes fainter than $M_{\text{F606}} \approx +2$, stars on the sub-RGB and near the MOTO (including any blue stragglers, if present in significant numbers in EXT8) with large photometric errors start scattering into the region of the optical CMD occupied by the HB, making the identification of potential extended BHB stars more ambiguous. While some stars with colours consistent with an extended BHB are visible down to $M_{\text{F606}} \approx +3$ in the optical CMD, it is not clear from Fig. 3 or Fig. 6 how many of them, if any, are actually HB stars. We therefore now turn to the analysis of the F300X images, which can help separate the populations in this part of the CMD.

In Fig. 12 (left column) we show the CMDs for stars detected in the F300X image while the right column shows the results of the artificial star tests. In addition to the HB and the brighter RGB stars, which have comparable F300X magnitudes, a small number of stars with magnitudes $M_{\text{F500X}} < 0$ are prominently visible above the HB. These stars are also visible in the optical CMDs, where they lie about a magnitude above the blue part of the HB (Fig. 11). These are possibly UV-bright post-HB or post-AGB stars (Zinn et al. 1972; Harris et al. 1983; Moehler et al. 2019; Bond 2021). Although some field contamination is possible in this region of the CMD (Fig. 3), the UV-bright candidates appear spatially associated with EXT8 with four of them located at radii $3'' < R < 7''$ and the fifth at about 15''. These stars are marked on the colour image in Fig. 1.

As in Fig. 3, the coloured circles in the right-hand panels of Fig. 12 show the completeness levels at various locations in the CMD according to the colour scale at the top. The black points in the right-hand panels are sources from the input artificial star catalogue that were recovered in the F300X photometry, while red points indicate sources that have no counterpart in the input catalogue and are mostly spurious detections. The errors depend on both the colours and magnitudes of the stars, but we have included error bars for stars with typical HB colours. As in Fig. 3, the error bars are based on the artificial star experiments. It is clear that the detection here strongly favours stars on the HB and we can now see that the HB in EXT8 does in fact extend at least to $M_{\text{F606}} \approx +3$. The completeness in F300X is about 50% at the level of the detected HB stars, but for colours bluer than $M_{\text{F300X}} - M_{\text{F606}} \approx -2$ the HB starts becoming fainter also in F300X and the detection efficiency rapidly drops. Deeper observations would be required to ascertain whether the HB extends even further and includes a more extreme blue tail as seen in some Galactic GCs (Fusi Pecci et al. 1993). A number of even fainter sources are visible in the lower left-hand corner of the $M_{\text{F300X}}$ versus $M_{\text{F300X}} - M_{\text{F606}}$ CMD, but comparison with the artificial cluster CMD shows that these are mostly spurious detections. This was further confirmed by an inspection of the spatial distribution of these sources, which showed them to be mostly random noise peaks, uniformly distributed across the detector. Adopting a more conservative detection threshold for the ALLFRAME photometry would lead to fewer such spurious detections, but would also decrease the detection efficiency for real sources.

As an alternative to the SMI index considered in Sect. 4.2.1 we can now consider the $\Delta(V - I)$ index, defined as the difference between the median $(V - I)$ colours of the HB stars and RGB stars at the level of the HB (Dotter et al. 2010). The median colour of the 50 RGB stars in the range $M_{\text{F606}} - \text{HB} = 0.25 < m_{\text{F606}} < m_{\text{F606}} = +0.25$ is $(M_{\text{F606}} - M_{\text{F814W}})_{\text{RGB}} = -0.33$. In addition to the 82 stars in the BHB box (Fig. 11), we counted 48 stars with $m_{\text{F606}} > 26.25$ and $M_{\text{F300X}} - M_{\text{F606}} < -1$ in Fig. 12. The median colour of the total sample of 130 HB stars
is \( M_{F606W} - M_{F814W} = -1.10 \), assuming that all HB stars with \( m_{F606W} > 26.25 \) have bluer colours than those in the BHB box. However, this still does not take into account the ~18 field contaminants in the BHB box, which are likely found mostly near the red boundary of the HB (Sect. 3.3). If we exclude the 18 reddest HB stars from the sample, the median HB colour shifts to \( M_{F606W} - M_{F814W} = -1.14 \). The median colour difference is then \( \Delta(F606W - F814W) = 0.81 \) mag or \( \Delta(V-I) = 0.95 \) mag. This is probably still an underestimate of \( \Delta(V-I) \) since we are likely missing a significant number of stars in the tail of the HB. Our estimate is therefore best expressed as a lower limit, \( \Delta(V-I) > 0.95 \), and the HB morphology as EXT8 might be as blue as the bluest HBs measured by Dotter et al. (2010).

To illustrate the effect of variations in helium abundance on the HB, Fig. 12 includes \( [\text{Fe/H}] = -3.2 \) BaSTI isochrones with \( Y = 0.247 \) and \( Y = 0.30 \) for an age of 13 Gyr. It is clear that the EXT8 HB extends beyond the bluest extent \( T_{\text{eff}} \approx 11,400 \) K of the \( Y = 0.247 \) isochrone. Stars with an increased helium content evolve faster and are therefore expected to arrive at the tip of the RGB (and then on the zero-age horizontal branch, ZAHB) with lower masses, and hence with less massive envelopes and correspondingly higher temperatures. For the BaSTI models in Fig. 12, the masses at the tip of the RGB are \( M = 0.705 M_\odot \) for \( Y = 0.247 \) and \( M = 0.638 M_\odot \) for \( Y = 0.30 \). The HB models do indeed extend further to the blue for \( Y = 0.30 \), reaching \( T_{\text{eff}} \approx 14,500 \) K, but still do not quite reach the most extreme HB stars that we detect in EXT8. A further increase in helium abundance would lead to still lower ZAHB masses and a correspondingly bluer HB, and the HB morphology might thus suggest some tentative evidence for the presence of relatively helium-rich stars in EXT8, perhaps related to the very low magnesium abundance inferred from the integrated-light spectrum.

It should be stressed, however, that many other second parameters besides helium abundance and age have been proposed (Fusi Pecci & Bellazzini 1997; Catelan 2009) and HB modelling in general is highly sensitive to the amount of mass loss along the RGB, which remains uncertain (Salaris et al. 2016). The effect of an increased mass loss can be mimicked by selecting an older isochrone: for an age of 17 Gyr, \( [\text{Fe/H}] = -3.2 \), and \( Y = 0.30 \), the ZAHB mass is \( M = 0.577 M_\odot \), which produces an HB that extends about as far as the faintest HB stars detected in EXT8 and reaches temperatures of \( T_{\text{eff}} \approx 18,000 \) K.

### 4.2.3. Variability

The \( F606W \) and \( F814W \) observations were obtained over two hours and five hours, respectively, and therefore do not fully sample typical RR Lyrae pulsation periods of 0.3–0.8 days (Oosterhoff 1939; Carney et al. 1992; Hoffman et al. 2021). Nevertheless, some variation on timescales of a few hours might be detectable, and as noted above the CMD does show a few stars located within the boundaries of the IS. Most of these are located near the blue edge and could be non-variables scattering into the IS from the part of the HB located just outside the IS. Conversely, some genuine variables might scatter out of the IS.

To look for evidence of variability among the HB stars, we plotted the rms variation of the \( F814W \) and \( F606W \) magnitudes measured on the individual \( \_flc \) images versus the mean errors reported by ALLFRAME for stars with \( -1.5 < M_{F606W} - M_{F814W} < -0.5 \) and with \( -0.5 < M_{F606W} < +1.0 \) (Fig. 13). We recall that the ALLFRAME errors on average reproduce the random errors on the measurements well (Sect. 3.4), so that a significantly larger rms variation than expected based on the errors might indicate a genuine change in brightness from one frame to another. The five stars located within the IS are marked with asterisks.

It should be stressed, however, that many other second variables might scatter out of the IS.

The outer star, which we refer to as RR Lyrae candidate A (RRC A), is located at a projected separation of about 19″ from the centre of EXT8. This is the star located furthest to the red in the CMD, inside the IS. RRC A is unaffected by crowding and we confirmed the variations in magnitude measured by ALLFRAME by remeasuring the star with the \texttt{imexamine} task in IRAF. The inner star, RRC B, is located about 3″/2 from the centre of EXT8. The light curves for the two candidates are shown in Fig. 14. Both stars show a relatively smooth variation in both filters, consistent with variability on timescales expected for RR Lyrae variables, although the coverage is too incomplete to determine periods. It is again worth recalling that even if these stars are genuine RR Lyrae variables, some field contamination is expected in this region of the CMD, so the RR Lyrae candidates could be field interlopers.

The right-hand panel in Fig. 13 shows several other stars with a relatively large rms variation in \( F814W \), \( \text{rms}(F814W) \gtrsim 0.15 \) mag, but smaller variations in \( F606W \). We found the distribution of these stars in the CMD to be similar to the general distribution of HB stars. Most of them tend to have relatively large errors (middle panel) and are located within a
radius of 5″ from the centre of EXT8 where crowding increases the photometric uncertainties.

4.3. A UV-bright source near the centre of EXT8

Inspection of the F300X image revealed a UV-bright source at a projected distance of about 0.3 or 1 pc west of the centre of EXT8. The source is quite distinct in the F300X image, but is much less conspicuous and suffers from blending with other nearby sources in F606W and F814W. To carry out photometry we first used the i-shape task (Larsen 1999) in IRAF to subtract a PSF-convolved model of EXT8 from the images. Figure 15 shows the F300X and F606W images of the central 2′ × 2′ of the cluster (left) and the residuals (right). We then obtained aperture photometry with the phot task in IRAF. In the F300X image we measured an apparent magnitude of $m_{F300X} = 21.87 \pm 0.04$ in an $r = 5$ pixels (0.′2) aperture, or $M_{F300X} = -3.19 \pm 0.04$ after applying an aperture correction of −0.24 mag. The source is thus more than 2 mag brighter in F300X than the other UV-bright stars discussed earlier (Sect. 4.2.2). In F606W it was difficult to obtain a meaningful measurement even on the model-subtracted image, owing to the very irregular background. For an $r = 5$ pixels aperture we found $m_{F606W} = 24.48 \pm 1.36$, while a smaller ($r = 3$ pixels) aperture yielded $m_{F606W} = 22.17 \pm 0.10$ and $M_{F606W} = 21.91 \pm 0.02$. Using the measurement in the smaller apertures we then find $M_{F300X} - M_{F606W} \approx -0.45 \pm 0.10$ as a rough estimate of the colour. However, the $m_{F606W}$ magnitude may be considered an upper limit, so the colour could be considerably bluer.

Comparing with stars on the HB, which have $M_{F300X} \approx +0.8$ and $L \approx 40 L_\odot$, the UV-bright source near the centre of EXT8 would have $L \approx 1600 L_\odot$. Again, the estimate of the bolometric luminosity $L$ is uncertain because of the large uncertainty on the colour. However, luminosities around 2000 $L_\odot$ are typical for post-AGB stars (de Boer 1985; Bond et al. 2021), such as Barnard 29 in M 13 (Dixon et al. 2019). Post-AGB stars evolve rapidly across the H–R diagram on timescales of $\sim 10^5$ yr and are therefore relatively rare, with Galactic GCs typically hosting between zero and two such stars (Moehler et al. 2019; Prabhu et al. 2021). Finding a single post-AGB star candidate in EXT8 is therefore not entirely unexpected. It may be noted that the detailed evolution of post-AGB stars remains uncertain, in part because of their relative rarity. It is unlikely that any large number of such stars remain undiscovered in the Milky Way GC system, but the observations presented here demonstrate that a search for post-AGB stars in M 31 GCs and other Local Group GCs should be quite feasible with space-based UV imaging.

4.4. Integrated properties and radial structure

To measure the integrated properties of EXT8 we used the photutils package (Bradley et al. 2020) in Python to carry out aperture photometry on the drizzle-combined images (the _drc files) in concentric apertures out to a radius of 12″, measuring...
We converted the WFC3 integrated photometry to standard Johnson–Cousins V and I magnitudes using the transformations in Harris (2018). This gives $V_0 = 15.28$, $M_V = -9.19$, and $(V - I)_0 = 0.75$, confirming that EXT8 is very blue for a GC as expected for its low spectroscopic metallicity. For comparison, Mackey et al. (2019b) found $(V - I)_0 = 0.79$ and $M_V = -9.28$, which makes EXT8 the bluest of the clusters included in their study. The difference between the $(V - I)_0$ colour reported by Mackey et al. (2019b) and that found here may be partly caused by uncertainties in the transformation from the WFC3 system to the Johnson–Cousins system, which is regarded as particularly difficult for the F606W filter and is likely uncertain by $\pm 0.1$ mag near $(V - I)_0 = 0.8$ (Harris 2018).

Half of the counts in the F606W image are contained within a half-light radius of $R_h = 0'092$ or about 3.5 pc, which is somewhat larger than the $R_h = 0'73 \pm 0'05$ (2.8 $\pm$ 0.2 pc) measured on a ground-based image by Larsen et al. (2020). There is a slight trend with wavelength, with $R_h = 0'90$ in the F300X image and $R_h = 0'95$ in F814W. The formal errors on the $R_h$ measurements due to counting statistics are small, $<0'01$. Systematic uncertainties are likely larger, but harder to estimate, and may be caused by uncertainties in the background subtraction, the centring of the apertures, and the behaviour of the profile beyond 12", as discussed further below. The relatively compact size of EXT8 follows the trend seen for other luminous GCs in the halo of M31 (Huxor et al. 2014).

We next used the kfit2d task in BA0LAB to fit a two-dimensional King (1966) profile to the drizzled F606W image of EXT8. The background was measured within the same 16’’–20’’ annulus used for the aperture photometry, and the fit was carried out within a radius of 12’’. The best fit was obtained for a model profile with a concentration parameter of $W_0 = 7.8$ and from this fit we found an ellipticity of $e = (1 - b/a) = 0.20 \pm 0.01$ (for minor/major axis ratio $b/a = 0.80 \pm 0.01$), a core radius measured along the major axis of $r_c = 0'238 \pm 0'001$, and a projected half-light radius of $R_h = 0'97 \pm 0'02$. The ellipticity is very similar to that measured from the ground-based image ($e = 0.19$; Larsen et al. 2020), while the half-light radius is again larger but agrees fairly well with the aperture photometry above.

Beyond a radius of about 12”, surface photometry becomes difficult due to the low surface brightness of the cluster and the two bright foreground stars. However, we used star counts to extend the profile out to about 25’’. Figure 16 shows the $\Delta_{\text{F606W-F814W}}$ offsets as a function of the semi-major axis $A$ of an ellipse drawn through each star. To calculate the semi-major axes we adopted the ellipticity and orientation of EXT8 from the two-dimensional King profile fits. To improve the statistics we included stars to a slightly fainter magnitude limit than in Sect. 4.1.3, $m_{\text{F606W}} = 25.0$, which is still expected to give a clean and complete sample of EXT8 RGB stars (Fig. 3). The sequence of RGB stars, centred on $\Delta_{\text{F606W-F814W}} = 0$, can be traced out to a semi-major axis of $A = 25''$, beyond which point very few stars with colours matching the EXT8 RGB stars are found.

In Fig. 17 we combine the data from Fig. 16 with surface photometry on the F606W image to show the surface brightness profile from the centre out to a semi-major axis of 25”’. For the surface photometry we used photutils to measure the flux in elliptical annuli. For the outermost regions we selected RGB stars with $|\Delta_{\text{F606W-F814W}}| < 0.07$ mag (red dashed lines in Fig. 16) and scaled the density profile obtained from the number counts such that it matched the surface photometry in the region of overlap (4’’ < A < 12’’).

Also shown in Fig. 17 are single-mass King (1966) and Wilson (1975) model profiles with concentration parameters of $W_0 = 7, 8, 9$. All model profiles have the same central surface brightness as the observed profile, $\mu_{\text{F606W}}(A = 0) = 15.39 \text{mag arcsec}^{-2}$ with a formal uncertainty less than 0.01 mag arcsec$^{-2}$. The profiles were calculated with the LITEMPy package (Gieles & Zocchi 2015) and have been scaled such that all model profiles have the same projected half-width at half maximum as the observed profile, HWHM = 0'23 measured along the semi-major axis. The HWHM is very similar to the core radius parameter $r_c$ (also sometimes called the King radius).
and has a similarly small formal uncertainty ($\approx 0.0^{+0.01}_{-0.01}$), but the conversion depends weakly on the concentration parameter. The core radii are between $0.24$ and $0.25$ for the profiles shown, similar to the value found from the two-dimensional fit. Because the radial profile analysis does not account for the PSF, the central surface brightness may be underestimated by $0.1$.

The various model profiles differ little in the central regions where they all fit the observed profile quite well. In the outer parts of the cluster, the King (1966)-type profiles do not provide good fits, while better fits are obtained for the more extended Wilson (1975) profiles. The best fit to the outermost parts is obtained for a $W_0 = 7.3$ Wilson profile, although such a profile has a slight deficit of light at intermediate radii (around $1''$) compared to the observations. There is also a hint that the observed profile drops somewhat more steeply beyond $A = 25''$, where extrapolation of the Wilson profile predicts an additional five $''$.

The exact age of EXT8 is not well constrained by the CMD, since it does not reach the MSTO and models predict a blue HB morphology even for GCs that are up to $\sim 2$ Gyr younger than a canonical ‘old’ GC at $[\text{Fe}/\text{H}] = -2.5$ (Catelan & de Freitas Pacheco 1993; Lee et al. 1994; Rey et al. 2001). We are not aware of specific predictions at the even lower metallicities. Although the HDF observations are most pronounced at $[\text{Fe}/\text{H}] < -1.5$, where the ages of the younger GCs appear to fit chemical enrichment models for potential progenitor dwarf galaxies in which the GCs may have resided before being accreted onto the halos of the larger galaxies (Kruissens et al. 2019; Forbes 2020). At metallicities $[\text{Fe}/\text{H}] < -2$ there is less evidence for any significant age spread among GCs (Dotter et al. 2011), a result that appears consistent with more general constraints on age-metallicity relations (Muratov & Gnedin 2010; Bellstedt et al. 2020; Horta et al. 2021). In all likelihood, EXT8 is at least as old as the oldest GCs observed in the Milky Way and elsewhere, although confirmation of this hypothesis must await deep photometry (for example with the James Webb Space Telescope) that can reach below the MSTO.

It might be argued that the HB morphology of EXT8 is unexpectedly blue, given that the bluest HBs tend to be found in clusters like M 10 and M 13 with $[\text{Fe}/\text{H}] = -1.5$ and not generally in the most metal-poor GCs (Renzini 1983; Sandage 1990; Dotter et al. 2010). Furthermore, the general tendency is for GCs at larger galactocentric distances to exhibit redder HB morphologies than those in the inner halos both in the Milky Way and M 31 (Lee et al. 1994; Perina et al. 2012). The very blue HB of EXT8 and its relatively large galactocentric distance may be contrary to this trend, although this must remain a tentative suggestion due to the uniquely low metallicity and corresponding lack of a suitable comparison sample.

Structurally, EXT8 resembles normal compact, luminous Galactic GCs, and it lacks the extended extra-tidal halos seen in some GCs that are suspected to be stripped nuclei (Da Costa 2015; Kuzma et al. 2018). It is not unusual for the more extended Wilson (1975)-type profiles to fit the outer regions of GCs better than the classical King (1966) profiles (McLaughlin & van der Marel 2005), so we find for EXT8 – both are relatively simple dynamical models that differ in their assumptions about the effect of a finite escape velocity on the distribution function (Gieles & Zocchi 2015). While the ellipticity of $\epsilon = 0.20$ indicates a noticeable degree of flattening, ellipticities up to 0.3 or more are observed among GCs in the Milky Way as well as other galaxies, with a tendency for relatively massive GCs to be more flattened (van den Bergh 1984; Harris et al. 2002; Gomez et al. 2006). Hence, EXT8 is not particularly exceptional in this respect either. We note that although rotation in a GC could in principle provide clues about its origins, there has not been a clear connection established in Galactic GCs between ellipticity and rotation (Kamann et al. 2018; Sollima et al. 2019).

It appears likely that the formation of EXT8 must be understood within the general context of GC formation scenarios. From the mass–metallicity relation in Choksi et al. (2018), a metallicity of $[\text{Fe}/\text{H}] = -2.9$ corresponds to a proto-galactic halo with a stellar mass of about $2 \times 10^8 M_\odot$ at a redshift of $z = 5$ and about $10^9 M_\odot$ at $z = 10$, which illustrates the fundamental difficulty of forming a GC this metal-poor and massive in a low-mass proto-galaxy, as envisioned in hierarchical galaxy formation scenarios (Choksi et al. 2018; Kruissens 2019). The difficulty is compounded if we consider that EXT8 probably did not form in isolation, although the inverse correlation between metallicity and GC specific frequency suggests that GCs might have formed more efficiently at lower metallicities (Harris & Harris 2002; Beasley et al. 2008; Lamers et al. 2017).

Of course, the actual mass–metallicity relation in the low mass, high-$z$ regime is poorly constrained from observations, and it remains unclear how much of a challenge a single object like EXT8 poses. It would clearly be of substantial interest to better
characterise the low-metallicity tail of the GC metallicity distribution and compare with simulations tailored to explore this part of parameter space. The existence of objects like EXT8 could inform models of galaxy enrichment in the extremely low metallicity regime (Wheeler et al. 2019; Agertz et al. 2020).

The fact that EXT8 stands out by its blue integrated colours suggests a route to identifying additional candidates, although colours are also sensitive to other parameters besides metallicity such as HB morphology, age, and variations in the MF (Vanderbeke et al. 2013). According to the Harris (1996) catalogue, the three metal-poorest Galactic GCs discussed in Sect. 4 have colours between $(V - I)_0 = 0.72$ (M 15) and $(V - I)_0 = 0.85$ (M 92) in spite of their very similar metallicities. It may here be noted, in particular, that M 15 is even bluer than EXT8, so a search for GC candidates with very blue colours is not guaranteed to turn up only extremely metal-poor objects. NGC 4147 has a similar integrated colour to EXT8, $(V - I)_0 = 0.76$, despite its significantly higher metallicity. Confirmation of candidate extremely metal-poor GCs will therefore require spectroscopic follow-up. The challenge here is that the metallic features become weak at such low metallicities, such that relatively high signal-to-noise spectra are needed for robust measurements.

6. Summary and conclusions

We have presented an analysis of HST observations of the extremely metal-poor GC EXT8 in M 31. The main conclusions are as follows:

- The CMD of EXT8 looks as expected for a metal-poor, old stellar population, with a steep RGB and a blue HB, and unambiguously confirms that EXT8 is an old, metal-poor GC.
- The colour of the upper EXT8 RGB is about 0.03–0.04 mag bluer in $M_{F606W} - M_{F814W}$ than the RGBs of the metal-poor Galactic GCs M 92 and M 30, but is similar in colour to the M 15 RGB. M 15 may, however, be more affected by uncertainties in the foreground reddening.
- The slope of the RGB is consistent with an extension of the relation between RGB slope and metallicity found by Sakari et al. (2015) to the lower metallicity of EXT8.
- The colour spread on the RGB is consistent with being caused entirely by observational uncertainties, although an intrinsic dispersion in $M_{F606W} - M_{F814W}$ of up to about 0.015 mag is allowed.
- Translating the upper limit on the colour dispersion to a limit on the allowed metallicity spread is very model dependent. According to MIST isochrones, the maximum colour spread corresponds to about 0.2 dex in metallicity, but when using BaSTI isochrones a much larger spread of at least 0.7 dex is allowed.
- The HB is located mainly on the blue side of the RR Lyrae gap and extends at least down to $M_{F606W} = +3$. A further extension of the HB would be below our detection limit and thus cannot be excluded. We quantify the HB morphology as $SMI \geq 0.9$ and $\Delta (V - I) > 0.95$, which is comparable to the bluest HBs observed in Galactic GCs.
- Two stars located in or near the instability strip have a larger rms scatter in $F606W$ and $F814W$ than expected from their mean photometric errors and may be considered RR Lyrae candidates.
- The CMD shows a few candidate UV-bright stars above the HB. These are the brightest stars in the $F300X$ versus $M_{F300X} - M_{F606W}$ CMD. An even brighter source is located near the centre and is a likely post-AGB star candidate.
- The surface brightness profile is well fitted by a Wilson (1975) model with a central surface brightness of $\mu_{F606W,0} = 15.2$ mag arcsec$^{-2}$ (corrected for foreground extinction), an ellipticity $\epsilon = 0.20$, and a core radius of $r_c = 0'25$ (about 0.93 pc) measured along the semi-major axis. There is no evidence of extra-tidal stars.
- From the RGB tip and the HB, we find that the distance modulus of EXT8 likely lies between $(m-M)_0 = 24.43 \pm 0.03$ and $24.57 \pm 0.05$. Depending on the distance of M 31 itself, the true three-dimensional distance between EXT8 and the centre of M 31 could be between the projected value (27 kpc) and up to about 70 kpc.

We conclude that EXT8 has properties consistent with it being a normal, albeit very metal-poor, GC. While its combination of relatively high mass and very low metallicity are challenging to explain in the context of GC formation theories operating within the hierarchical galaxy assembly paradigm, a more systematic search for similar objects would be highly desirable.

The results of this work reaffirm the dual roles that GCs can play in bridging stellar and (extra-)galactic astrophysics. Finding more objects like EXT8 would not only help constrain galaxy formation theories, but would also provide valuable input for the calibration of stellar model properties such as the RGB colour, HB morphology, and post-AGB evolution at low metallicities.

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# Appendix A: Photometry

## Table A.1. Photometry in $F606W$ and $F814W$.

| ID | X     | Y     | $F606W$ | $F814W$ |
|----|-------|-------|---------|---------|
|    |       |       | mag     | error   | rms    | mag     | error   | rms    |
| 1  | 52.959| 52.579| 26.056  | 0.120   | 0.152  | 26.570  | 0.117   | 0.126  |
| 5  | 4047.204| 52.729| 26.053  | 0.121   | 0.125  | 26.559  | 0.114   | 0.202  |
| 24 | 3961.471| 52.602| 27.827  | 0.138   | 0.095  | 28.231  | 0.165   | 0.233  |
| 36 | 3945.614| 53.102| 27.980  | 0.099   | 0.272  | 27.952  | 0.124   | 0.259  |
| 40 | 623.995| 53.526| 27.561  | 0.079   | 0.182  | 27.610  | 0.079   | 0.229  |

**Notes.** Only the first few rows are reproduced here; the full table is available at the CDS. $X$ and $Y$ are the coordinates on the WFC3 UVIS2 $F606W$ image.

## Table A.2. Photometry on drizzled images in $F300X$, $F606W$, and $F814W$.

| ID | X     | Y     | $F300X$ | $F606W$ | $F814W$ |
|----|-------|-------|---------|---------|---------|
|    |       |       | mag     | error   | chi     | mag     | error   | chi     |
| 4  | 23.631| 3.019 | 26.771  | 0.629   | 3.53    | 29.283  | 0.702   | 2.01    |
| 5  | 30.314| 3.110 | 26.903  | 0.340   | 1.71    | 26.444  | 0.331   | 6.21    |
| 7  | 38.383| 2.947 | 27.393  | 0.675   | 2.31    | 28.756  | 0.361   | 1.71    |
| 15 | 2.689 | 8.221 | 22.539  | 0.249   | 16.75   | 30.129  | 1.386   | 1.74    |
| 17 | 17.889| 10.396| 27.261  | 0.514   | 1.91    | 28.322  | 0.217   | 1.36    |

**Notes.** Only the first few rows are reproduced here; the full table is available at the CDS. $X$ and $Y$ are the coordinates on the drizzled WFC3 $F300X$ image.