Probing IGM accretion onto faint Lyα emitters at $z \sim 2.8$

Fakhri S. Zahedy$^{1,2,\star}$, Michael Rauch$^2$, Hsiao-Wen Chen$^{1,3}$, Robert F. Carswell$^4$, Brian Stalder$^5$, Antony A. Stark$^6$

$^1$Department of Astronomy & Astrophysics, The University of Chicago, Chicago, IL 60637, USA
$^2$The Observatories of the Carnegie Institution for Science, 813 Santa Barbara Street, Pasadena, CA 91101, USA
$^3$Kavli Institute for Cosmological Physics, The University of Chicago, Chicago, IL 60637, USA
$^4$Institute of Astronomy, University of Cambridge, Cambridge CB3 0HA, UK
$^5$LSST, 950 North Cherry Avenue, Tucson, AZ 85719, USA
$^6$Center for Astrophysics | Harvard & Smithsonian, Cambridge, MA 02138, USA

26 March 2019

ABSTRACT

Observing the signature of accretion from the intergalactic medium (IGM) onto galaxies at $z \sim 3$ requires the detection of faint ($L \ll L^*$) galaxies embedded in a filamentary matrix of low-density ($\rho < 100 \rho$), metal-poor gas ($Z \sim 10^{-2.5} Z_\odot$) coherent over hundreds of kpc. We study the gaseous environment of three Lyα emitters (LAEs) at $z = 2.7 - 2.8$, found to be aligned in projection with a background QSO over $\sim 250$ kpc along the slit of a long-slit spectrum. The lack of detection of the LAEs in deep continuum images and the low inferred Lyα luminosities show the LAEs to be intrinsically faint, low-mass galaxies ($L \ll 0.1 L^*$, $M_\text{star} \ll 0.1 M^*$). An echelle spectrum of the QSO reveals strong Lyα absorption within $\pm 200$ km s$^{-1}$ from the LAEs. Our absorption line analysis leads to H$\text{I}$ column densities in the range of $\log N(\text{H}\text{I})/\text{cm}^{-2} = 16 - 18$. Associated absorption from ionic metal species C$\text{IV}$ and Si$\text{IV}$ constrains the gas metallicities to $\sim 0.01$ solar if the gas is optically thin, and possibly as low as $\sim 0.001$ solar if the gas is optically thick, assuming photoionization equilibrium. While the inferred metallicities are at least a factor of ten lower than expected metallicities in the interstellar medium (ISM) of these LAEs, they are consistent with the observed chemical enrichment level in the IGM at the same epoch. Total metal abundances and kinematic arguments suggest that these faint galaxies have not been able to affect the properties of their surrounding gas. The projected spatial alignment of the LAEs, together with the kinematic quiescence and correspondence between the LAEs and absorbing gas in velocity space suggests that these observations probe a possible filamentary structure. Taken together with the blue-dominant Lyα emission line profile of one of the objects, the evidence suggests that the absorbing gas is part of an accretion stream of low-metallicity gas in the IGM.

Key words: galaxies:haloes – galaxies: high-redshift – quasars: absorption lines – galaxies: abundances

1 INTRODUCTION

The presence of heavy elements in the intergalactic medium (IGM) can be used to probe the relationship between galaxies and the IGM. Absorption lines caused by common metals in the spectra of background QSOs, when observed in sightlines close to foreground galaxies, and, occasionally, supplemented by observations of nebular emission from the gas, can shed light on how galaxies accrete gas and smaller building blocks from the ambient IGM to fuel star formation, and subsequently eject metals and radiation to their environment.

Adding metals ab-initio to a cosmological hydro-simulation of the IGM, Haehnelt et al. (1996) showed that typical QSO metal absorbers can arise during the inflow of pre-enriched gas ($Z \lesssim -2.5 Z_\odot$) into $z \sim 3$ galactic potential wells of progenitors ($M_\text{baryon} \sim 10^{9-10} M_\odot$) of galaxies like the Milky Way. The close agreement between the properties of the simulated and observed $z \sim 3$ metal absorption lines, even in the absence of any local metal enrichment or galactic

$\star$ E-mail: fsz@uchicago.edu

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feedback (Rauch et al. 1997), suggest that metal absorption systems (e.g., C iv) may provide our earliest observational glimpses of the accretion of gas onto high redshift galaxies.

The early-enrichment or infall scenario for QSO metal absorbers can be contrasted with a more local origin for the heavy metals, where feedback from recent star formation activity enriches the gaseous environment of a galaxy with metals (the circumgalactic medium, CGM; e.g. Adelberger et al. 2003; Steidel et al. 2010). However, the existence of long-range outflows able to explain the widespread metal enrichment through local feedback has been questioned (e.g., Kawata & Rauch 2007; Gauthier & Chen 2012). Furthermore, the observed low levels of turbulence and spatial small-scale structure in the IGM seen in multiple lines-of-sight to lensed QSOs broadly support the idea that most metal absorbers are seen in stages of inflow as part of the general IGM, unaffected by recent feedback (e.g., Rauch et al. 2001a,b). A study of the CGM of Lyman break galaxies has confirmed that the redshift distortions in the velocity field of H i, Si iv and C iv absorption at $z = 2 - 3$ are in fact due to infall, not outflows (Turner et al. 2017).

The formation of QSO metal absorbers from infalling IGM gas requires an earlier epoch of widespread metal enrichment in the IGM, the existence of which has been established by detections of ionic metal species at high redshifts currently up to $z \sim 6.8$ (e.g., Songaila 2001; Schaye et al. 2003; Simcoe et al. 2004, 2011; Ryan-Weber et al. 2009; Becker et al. 2011; Chen et al. 2017; Cooper et al. 2019). Theoretical and empirical arguments also support the idea that the IGM was likely polluted by more widely distributed, low-mass galaxies at earlier times. Gas escapes the shallow potential well of low-mass objects more easily and high-redshift dwarf galaxies occur at high number densities, favoring a widespread pollution of the IGM at early times, aided by the Hubble expansion (e.g., Scannapieco 2005; Pocrani & Madau 2005; Wiersma et al. 2010).

Recent advances in cosmological hydrodynamic simulations have refined predictions for galactic accretion and emphasized the importance of small scale structure embodied by cold accretion streams (e.g., Kereˇ s et al. 2005; Goerdt et al. 2010; Faucher-Giguere & Kereˇ s 2011; Fumagalli et al. 2011; Rosdahl & Blaizot 2012; Van de Voort et al. 2012; Shen et al. 2013). Comparisons between observations and the expected accretion features in these cosmological simulations need to rely on statistical arguments, focusing on a subset of properties like column density or metallicity distributions (e.g., Cooper et al. 2015; Hafen et al. 2017).

In contrast, direct detections of accretion onto individual high-redshift galaxies have remained rare, as they must rely on multiple strands of evidence, essentially requiring two- or three-dimensional information that can only be provided by simultaneous observations of gas and stars in emission. Evidence for gaseous accretion onto galaxies may include kinematic signatures of gas streams attached to and/or moving towards a galaxy, gaseous filaments converging on starburst regions fueled by the inflow, and tidal tails from the simultaneous infall of small halos. One candidate system for filamentary infall onto a high-redshift galaxy was reported by Rauch et al. (2011, 2016). That particular system owes its discovery to fluorescent radiation from gaseous features which have exposed to ionizing photons escaping from the starbursting region. However, such a situation may be atypical of galaxies undergoing the cold accretion process, most of which will have too low star-formation rates and presumably masses required to illuminate their gaseous environment.

Observing more typical forms of galactic accretion requires the ability to simultaneously detect fainter galaxies and fainter nebular emission features. In this paper we will study such a putative accretion environment, by paying a closer look at a grouping of $z \sim 2.8$ Lyα emitters (LAEs) previously discovered through a blind spectroscopic long-slit
search (Rauch et al. 2008, hereafter R08). As shown in Figure 1, this grouping is unique because the LAEs are coincident in redshift with a number of Lyα absorption complexes seen in the sightline of a nearby background QSO, DMS 2139–0405. The underlying galaxies of the LAEs are faint, with low inferred star formation rates of SFR \(< 1 \, M_\odot \, yr^{-1}\). In particular, one of the gas complexes, Complex A, occurs in a structure spanned by a pair of Lyα-emitting galaxies which are situated at projected distances \(d = 87\) and 161 kpc on opposite sides of the QSO sightline. One of the LAEs shows a relatively rare Lyα emission profile with a dominant blue peak (R08), which is commonly associated with infalling gas (e.g., Verhamme et al. 2006).

The positional information about the LAEs derived from the two-dimensional low resolution spectra of the field, and the high-resolution absorption spectra of neutral hydrogen and metal species along the QSO sightline, allow us to study the physical association between the galaxies and the gas, and to test the predictions of the infall scenario for metal absorption systems against the observed gas column density, chemical abundance, ionization state, and kinematics. The low luminosities and likely low masses of the LAEs also enable us to examine the validity of the late-time outflow scenario and the CGM paradigm in a more typical environment of the cosmic web far from massive starburst galaxies.

This paper is organized as follows. Section 2 presents the observational dataset and corresponding data reduction. We describe the properties of the foreground LAEs in Section 3. In Section 4, we present the absorption-line measurements, describe the absorption line complexes near the LAEs, and investigate the ionization and chemical properties of the gas. Finally, we discuss the interpretations of our findings in Section 5 and summarize our study in Section 6. Throughout this paper, we adopt a \(\Lambda\) cosmology of \(\Omega_M = 0.3\) and \(\Omega_\Lambda = 0.7\), with a Hubble constant of \(H_0 = 70\) km s\(^{-1}\) Mpc\(^{-1}\). All magnitudes are reported in the AB system.

2 OBSERVATIONS

To characterize the properties of the absorbing gas and investigate its connection to the unique grouping of \(z \sim 2.8\) LAEs, we carried out high-resolution echelle spectroscopy of the background QSO DMS 2139–0405 (\(z_{\text{QSO}} = 3.32\)). In addition, we obtained new optical and near-infrared broadband imaging data of the field around the QSO in order to search for continuum counterparts to the LAEs. Here we describe the observations and data reduction procedures.

2.1 QSO echelle spectroscopy

We obtained high-resolution echelle spectra of QSO DMS 2139–0405 (right ascension 21\(^{\text{h}}41^{\text{m}}39.1^{\text{s}},\) declination \(-03\degree 51\arcmin 42.57\arcsec\) in the J2000 epoch) with the MIKE echelle spectrograph (Bernstein et al. 2003) on the Magellan Clay Telescope over two nights in 2016 September. A 1-arcsec slit and 3 \(\times\) 3 binning during readout was chosen for the observations, delivering a spectral resolution of FWHM \(\approx 10 – 12\) km s\(^{-1}\) over a continuous wavelength coverage from 3350 Å to 9300 Å. The mean seeing over the observing period was 0.5 – 0.7 arcsec. The total integration time of the MIKE observations was 26400 s, comprising eight individual exposures of equal duration. We reduced the MIKE spectra using an idl-based custom data reduction software. For each individual exposure, the software performed optimal extraction of the QSO spectrum in each echelle order, using a Gaussian weighting scheme that matched the observed QSO spatial profile. Observations of a spectrophotometric standard star taken during the same night as the observations were utilized to determine the response function, which was then used to perform relative flux calibrations of the QSO spectrum. The individual echelle orders were combined for each exposure, and the different exposures were added to form a single spectrum, which was then continuum normalized using a low-order polynomial function that excluded regions of strong absorption. The final reduced spectrum of DMS 2139–0405 is characterized by a signal-to-noise ratio of S/N \(\approx 10 – 15\) per resolution element at \(\lambda > 4000\) Å.

2.2 Imaging observations

Optical imaging data of the field around DMS 2139–0405 were obtained using the Panoramic Imager for Southern Cosmology Observations (PISCO; Stalder et al. 2014) on the Magellan Clay Telescope in 2016 October. PISCO is a multi-band imager which provides simultaneous broadband coverage in \(g, r, i,\) and \(z\)-bands. The observations consisted of a series of exposures 300 s in length, with a total integration time of 3300 s taken under a mean seeing of 0.9 – 1.0 arcsec. The raw PISCO data were reduced using standard idl-based routines. The photometric zero points of the PISCO imaging data were determined using field stars observed in the Sloan Digital Sky Survey (SDSS; York et al. 2000). We supplement our PISCO imaging data with archival \(V\)- and \(I\)-band images of the field taken with LRIS (Oke et al. 1995) on the Keck-I Telescope, which were previously described in R08. The LRIS \(V\)- and \(I\)-band images are characterized by a mean seeing of 0.7 – 0.9 arcsec.

In addition to the optical images, we obtained near-infrared (NIR) \(K_s\)-band observations of the field with FourStar (Persson et al. 2008) on the Magellan Baade Telescope in 2016 August. The field was observed for a total of 14120 s, under excellent seeing conditions of 0.5 – 0.8 arcsec. To prevent saturation due to high sky and thermal backgrounds in the image, the observations consisted of a sequence of short 20 s exposures with a 10-arcsec random dither applied between successive exposures. The FourStar data were reduced using an idl-based routine which per-
formed sky subtraction and flat-fielding using super sky frames created on a rolling basis. We used field stars observed with the Two Micron All Sky Survey (2MASS; Skrutskie 2006) to determine the photometric zero point of the $K_s$-band image.

We summarize the available imaging data of the field around QSO DMS 2139–0405 in Table 1, where we also show the mean 5σ limiting magnitudes measured over a 1 arcsec sky aperture for each bandpass. The limiting magnitudes range from 25.1 mag for NIR $K_s$ band to deeper than 27 mag for the optical $g$, $V$, and $r$-bands. In Figure 2, we present a false-color composite image of the field around QSO DMS 2139–0405. The composite image was created using the $V$, $I$, and $K_s$-band images for blue, green, and red, respectively.

3 FOREGROUND LAEs AT $d < 300$ kpc FROM DMS 2139–0405

To study the gaseous environment of high-redshift LAEs, we focus on foreground LAEs within $d < 300$ kpc from the QSO sightline. Two of these objects, LAE A1 at peak redshift $z_{\text{peak}} = 2.7659$ and LAE A2 at $z_{\text{peak}} = 2.7713$, are closely situated in redshift space with their Ly$\alpha$ emission peaks separated by merely $\Delta v \approx 400$ km s$^{-1}$. In addition, both LAEs are situated on opposite sides of the QSO at a projected separation of 248 kpc, with the QSO sightline probing the LAE A1 at $d = 87$ kpc and LAE A2 at $d = 161$ kpc (top-left panel of Figure 3). A third emitter, LAE B1, occurs at $z_{\text{peak}} = 2.7483$, and is probed by the background QSO at $d = 187$ kpc (top-left panel of Figure 4).

Because the shape of Ly$\alpha$ emission line is highly sensitive to details of the radiative transfer (e.g., Verhamme et al. 2006), it is likely that the systemic redshifts of the LAEs, $z_{\text{sys}}$, are different from their measured Ly$\alpha$ peak redshifts, $z_{\text{peak}}$. Previous studies found that a large majority of Ly$\alpha$-selected galaxies at $z \sim 2$–3 exhibit asymmetric Ly$\alpha$ emission line profiles, with a main peak that is typically redshifted from the systemic redshift by greater than 100 km s$^{-1}$ (e.g., Erb et al. 2014; Hashimoto et al. 2015; Trainor et al. 2015). Using a sample of faint LAEs, Hashimoto et al. (2015) found a mean offset of $\Delta v_{\text{red}} = +174$ km s$^{-1}$ for the main red peak relative to systemic. To estimate the systemic redshifts of the LAEs, we adopt the mean offset from Hashimoto et al. (2015) and apply it to the peak redshifts of LAEs A2 and B1. For LAE A1, which exhibits a dominant blue peak in its Ly$\alpha$ emission profile (R08), we estimate its systemic redshift by adopting an offset of $\Delta v_{\text{blue}} = -316$ km s$^{-1}$, which is the mean offset Hashimoto et al. (2015) found for blueshifted Ly$\alpha$ emission line peaks.

The observed Ly$\alpha$ line fluxes of the three LAEs are $F_{\text{Ly} \alpha} = (2.7 \pm 0.4) \times 10^{-18}, (3.3 \pm 0.6) \times 10^{-18}$, and $(3.5 \pm 0.5) \times 10^{-18}$ erg s$^{-1}$ cm$^{-2}$ for LAEs A1, A2, and B1 (R08). These observed line fluxes correspond to Ly$\alpha$ luminosities of $L_{\text{Ly} \alpha} = (1.7 \pm 2.1) \times 10^{31}$ erg s$^{-1}$. The inferred Ly$\alpha$ luminosities of the three LAEs are a factor of 20–30 lower than the characteristic luminosity of the LAE luminosity function at $z \sim 3$ (e.g., Cassata et al. 2011; Konno et al. 2016; Drake et al. 2017), which indicate that these LAEs are intrinsically faint compared to the general LAE population at the same epoch.

The available multi-band imaging data of the field around DMS 2139–0405 allows us to search for continuum light counterparts of these faint LAEs. We infer the positions of the LAEs on the sky based on the reported long-slit position and orientation for the R08 observations as well as the spatial position of the Ly$\alpha$ emissions in the long-slit spectrum from R08. For each bandpass, we then measure the total flux contained within a 3-arcsec diameter aperture centered at the inferred position of each LAE. The choice of aperture size is guided by the width of the long slit used.

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1 This conclusion would stand even if the observations were subject to significant slit losses of ~50% (e.g., R08; Cassata et al. 2011).
in R08. We find that none of the three LAEs is detected in continuum light at 3σ or higher significance level, in all broadband filters. The lack of significant detection in any broadband image persists if a smaller (2 arcsec) or larger (5 arcsec) aperture is adopted instead.

The lack of detection of continuum light from all three LAEs (r > 26.5, Ks > 24.5) indicates that these LAEs are faint and low-mass galaxies. At z = 2.7–2.8, the upper limit on r-band flux corresponds to a rest-frame FUV absolute magnitude of $M(1700 \text{ Å}) \approx 18.7$, which is equivalent to $L \lesssim 0.1 L^*$ (e.g., Reddy et al. 2008). Using this limit on FUV absolute magnitude, we constrain the total stellar masses of the LAEs to $M_\text{star} \lesssim 10^9 M_\odot$, based on the inferred mean mass-to-light ratio and stellar mass function as this epoch (e.g., Reddy et al. 2012; Marchesini et al. 2009). The inferred low masses of the three LAEs are consistent with the low mean dark matter halo mass ($M_\text{h} \sim 10^{11} M_\odot$) of z ~ 3 LAEs inferred from clustering analyses (e.g., Bipley et al. 2016).

Next, we estimate the star-formation rate (SFR) of the LAEs. At z ≈ 2.8, the effective wavelength of the i-band filter corresponds to a rest-frame wavelength of $\approx 2050 \text{ Å}$ in the near-ultraviolet (NUV). By converting the i-band flux limits of the LAEs to their corresponding upper limits on specific luminosity in the NUV, and then adopting the local NUV star-formation rate relation from Kennicutt & Evans (2012, equation 12), we estimate upper limits on unobscured SFR of SFR$_{\text{UV}} \lesssim 3.0 M_\odot$ yr$^{-1}$ for the three LAEs.

An independent estimate of the star-formation rate of the LAEs can also be obtained by assuming that the observed Lyα emission is powered by underlying star formation activity in the galaxies. If we also assume case-B recombination (Brocklehurst 1971), the Kennicutt & Evans (2012) SFR relation for Hα can be converted into one for Lyα, following $\text{SFR}_{\text{Ly} \alpha} (M_\odot \text{ yr}^{-1}) = 6.2 \times 10^{-43} L_{\text{Ly} \alpha} (\text{erg s}^{-1})$. Applying this relation to the observed Lyα luminosities gives unobscured SFRs of $\text{SFR}_{\text{Ly} \alpha} \approx 0.1 M_\odot \text{ yr}^{-1}$ for the three LAEs.

In Table 2, we summarize the photometric and spectroscopic properties of LAEs A1, A2, and B1. For each object, we report its angular offset from the QSO, observed peak Lyα redshift, estimated systemic redshift, total Lyα line flux, projected distance from the QSO sightline, 3σ upper limits on the broadband fluxes, and the inferred SFRs. For completeness, we also include the same information in Table 2 for LAEs 1, 2, and 3, which are located at z ~ 3.3 or less than a few $1000 \text{ km s}^{-1}$ in velocity separation from the QSO.

### 4 ANALYSIS AND RESULTS

The high-resolution echelle spectrum of DMS 2139–0405 enables a detailed study of gas in proximity to the three foreground LAEs. As previously mentioned in §1, a number of Lyα absorption complexes at z = 2.7–2.8 are found in the low-resolution spectrum of the QSO (Figure 1). In particular, Complex A and Complex B are coincident in redshift space with the three foreground LAEs situated at $d < 300 \text{ kpc}$ (see Figure 1 and §3). In this section, we describe the observed gas properties in the vicinity of these $z = 2.7–2.8$ LAEs, based on our analysis of the high-resolution echelle spectrum of DMS 2139–0405.

### 4.1 Absorption Analysis

At the LAE redshifts, our MIKE spectrum of DMS 2139–0405 provides coverage of the H I Lyα $\lambda 1215$ line and prominent ionized metal transitions C IV $\lambda 1548, 1550$ and SiIV $\lambda 1393, 1402$ doublets. To characterize the ab-
sorbers, we first measure the total rest-frame equivalent widths for transitions $\text{H}\alpha$, $\text{Ly}\alpha$, $\text{C}\text{II}\lambda 1548, 1550$, and $\text{Si}\text{IV}\lambda 1393, 1402$ absorption lines. When a given ionic transition is not detected, we use the error array to estimate the $2\sigma$ upper limit on the absorption equivalent width, integrated over a spectral window that is twice the FWHM of the corresponding $\text{H}\alpha$ line. The equivalent width measurements are summarized in Table 3, where for each transition we present the absorption redshift $z_{\text{abs}}$, the rest-frame equivalent width $W_{\text{r}}$, and the velocity interval over which the absorption equivalent width is integrated.

To further investigate the gas properties, we perform a Voigt profile fitting analysis using the VPFIT package (Carnell & Webb 2014), in order to constrain the centroid redshift, $z_{\text{c}}$, gas column density, $\log N_{\text{c}}$, and Doppler parameter, $b_{\text{c}}$, of individual absorbing components. We report the results of the Voigt profile fitting in Table 4, and present the best-fit model absorption profiles for both absorption complexes along with the data in Figure 3 and Figure 4. Uncertainties in Table 4 represent the estimated 1σ (68%) confidence level from the fitting analysis. When no $\text{CIV}$ or $\text{SiIV}$ absorption is detected, we estimate the $2\sigma$ upper limit on the absorption equivalent width of the strongest available transition of the species, which is calculated over a velocity window twice the full-width-at-half-maximum (FWHM) of the corresponding $\text{Ly}\alpha$ component. Then, we calculate the corresponding $2\sigma$ upper limit on that component’s column density assuming the gas is optically thin. For saturated $\text{Ly}\alpha$ components, we also estimate the range of allowed $\log N_{\text{c}}$ and $b_{\text{c}}$ at the 95 percent confidence level, based on a likelihood contour constructed from a grid of $\chi^2$ values from comparing Voigt profile models and observations.

### 4.2 Gaseous environment of three $z = 2.7 - 2.8$ LAEs

As discussed in §3, three LAEs occur at $d < 300$ kpc from the background QSO sightline (Figure 1). We now describe two $\text{Ly}α$ absorption complexes which are coincident with these emitters in velocity space (see Figure 1). Each absorption system is found within a line-of-sight velocity difference of $|\Delta v| < 500$ km s$^{-1}$ from the LAEs.

Complex A is found near the redshifts of LAE A1 and LAE A2 (see absorption panels in the bottom of Figure 3). Our MIKE echelle spectrum of DMS 2139−0405 reveals strong $\text{H}\alpha$ absorption with a total rest-frame equivalent width of $W_{\text{r}}(1215) = 1.24 \pm 0.02$ Å. Corresponding strong $\text{CII}$ and modest $\text{SiIV}$ metal absorptions are also detected. The metal absorption profiles exhibit a multi-component structure that kinematically flank the systematic redshifts of the two LAEs. By fitting a single absorption component to the observed $\text{Ly}α$ absorption, our Voigt profile fitting procedure returns a best-fit $\log N(\text{H}α)/\text{cm}^{-2} = 17.1 \pm 0.3$ for this system. The large uncertainty in $N(\text{H}α)$ is understood as due to the saturated $\text{Ly}α$ MIKE profile and the fact that no higher-order $\text{H}\alpha$ transition is available in the MIKE data. Based on $\chi^2$ values from the fitting analysis, we estimate that the range of allowed $\text{H}\alpha$ column density is $\log N(\text{H}α)/\text{cm}^{-2} = 16.5 - 17.7$ at the 95 percent confidence level, with a corresponding Doppler parameter range of $b = 52 - 64$ km s$^{-1}$. The range of allowed total $N(\text{H}α)$ does not change significantly if we instead impose a multi-

![Figure 3. Top-left: White-light image of the field around QSO DMS 2139−0405. The slit orientation of the R08 long-slit observations is shown in parallel red lines. The positions of foreground LAEs A1 and A2 are indicated in yellow squares. The LAEs are situated on opposite sides of the QSO, with LAE A1 at $d = 87$ kpc and LAE A2 at $d = 161$ kpc. Top-right: A cutout of the two-dimensional long-slit spectrum, centered in spectral (horizontal) and spatial (vertical) directions at the $\text{Ly}α$ absorption from absorption complex A. The red (blue) arrow points to the location of $\text{Ly}α$ emission from LAE A1 (LAE A2). Bottom: Continuum-normalized absorption profiles for various ionic transitions detected in absorption complex A. Zero velocity corresponds to the redshift of the $\text{Ly}α$ absorption, $z = 2.7691$. The absorption spectra and corresponding 1σ error arrays are presented in black and cyan histograms. The dashed blue (red) line marks the adopted systemic redshift of LAE A1 (LAE A2) (see §3). The best-fitting Voigt profile models are presented in red curves, with magenta tick marks indicating the centroid of individual components. For the saturated $\text{Ly}α$ line, we estimate that the range of allowed $\log N(\text{H}α)/\text{cm}^{-2} = 16.5 - 17.7$ at 95 percent confidence level, with a corresponding Doppler parameter range of $b = 52 - 64$ km s$^{-1}$.](https://example.com/figure3)
Table 4. Voigt profile fitting results for z \sim 2.8 absorption complexes associated with LAEs

| Species | z_c | log N_c/cm^-2 | b_c (km/s^-1) |
|---------|-----|---------------|---------------|
| Complex A at z = 2.7691 |
| SiIV | 2.76883 ± 0.00008 | 12.11 ± 0.24 | 7.9 ± 8.6 |
| CIV | 2.76893 ± 0.00001 | 13.70 ± 0.04 | 12.5 ± 1.6 |
| HI | 2.76914 ± 0.00001 | 17.1 ± 0.3 | 57 ± 3 |
| SiIV | 2.76926 ± 0.00004 | 12.11 ± 0.21 | 4.7 ± 8.9 |
| CIV | 2.76933 ± 0.00002 | 13.10 ± 0.24 | 3.4 ± 4.6 |
| CIV | 2.76959 ± 0.00001 | 13.81 ± 0.13 | 63 ± 1.6 |
| SiIV | 2.76962 ± 0.00002 | 12.47 ± 0.11 | 6.5 ± 4.3 |
| CIV | 2.77002 ± 0.00002 | 13.53 ± 0.05 | 12.8 ± 2.2 |
| SiIV | 2.77014 ± 0.00005 | 12.15 ± 0.20 | 6.0 ± 7.3 |
| Complex B at z = 2.7440 |
| HI | 2.7418 ± 0.0002 | 12.98 ± 0.29 | 25.9 ± 16.6 |
| CIV | ... | < 12.8 | ... |
| SiIV | ... | < 12.5 | ... |
| CIV | 2.74398 ± 0.00006 | 13.53 ± 0.07 | 40.4 ± 7.2 |
| HI | 2.74400 ± 0.00002 | 17.3 ± 0.5 | 52 ± 2 |
| SiIV | ... | < 12.7 | ... |
| HI | 2.74755 ± 0.00003 | 13.36 ± 0.06 | 21.8 ± 3.0 |
| CIV | ... | < 12.8 | ... |
| SiIV | ... | < 12.5 | ... |
| HI | 2.74915 ± 0.00002 | 14.56 ± 0.06 | 51.5 ± 3.1 |
| CIV | ... | < 12.9 | ... |
| SiIV | ... | < 12.7 | ... |
| HI | 2.75127 ± 0.00004 | 13.37 ± 0.11 | 36.7 ± 6.5 |
| CIV | ... | < 12.9 | ... |
| SiIV | ... | < 12.3 | ... |

component HI profile that match the kinematic structure of the observed CIV complex.

The CIV and SiIV absorption profiles in Complex A can each be separated into a minimum of four discrete components which are spread over almost 100 km s^{-1} in line-of-sight velocity. We measure a total ion column density of \log N(CIV)/cm^{-2} = 14.21 ± 0.06 for CIV and \log N(SiIV)/cm^{-2} = 12.84 ± 0.09 for SiIV, integrated over all components.

Complex B is found at the redshift of LAE B1. In contrast to the more kinematically compact absorption profile of Complex A, the Lyα profile of Complex B consists of multiple components spread over nearly 900 km s^{-1} in line-of-sight velocity (see the bottom panels of Figure 4). The integrated rest-frame equivalent width of the Lyα absorption is W_{Lyα}(1215) = 1.99 ± 0.03 Å. Our Voigt profile analysis shows that the total N(HI) in Complex B is dominated by a saturated Lyα component centered at Δv = −170 km s^{-1} blueward of the redshift of LAE B1. While the fitting procedure yields a best-fit \log N(HI)/cm^{-2} = 17.3 ± 0.5 for this component, a full \chi^2 analysis indicates that the range of allowed column density is \log N(HI)/cm^{-2} = 16.5 − 17.8 at the 95 percent confidence level, with a corresponding Doppler linewidth of \theta = 48 − 61 km s^{-1}. In addition to H1, we detect modest CIV absorption in this complex, at a location that matches the strongest Lyα component in velocity space. The best-fitting Voigt profile model for the CIV absorption is given by a single-component with
Figure 5. Example of predictions from CLOUDY photoionization calculations. The top two panels show the ionic column density ratio $\log N(C^{\text{iv}})/N(Si^{\text{iv}})$ as a function of ionization parameter $U$. The calculations are performed for an optically thin gas ($\log N(H^\text{i})/cm^{-2} = 16.5$) on the left and optically thick gas ($\log N(H^\text{i})/cm^{-2} = 17.7$) on the right, matching the allowed minimum and maximum $N(H^\text{i})$ of the absorption complex A (see § 4.1). In the bottom two panels, the ionization fractions of $C^{+3}$ and $Si^{+3}$ relative to $H^0$ are plotted versus $U$ for both optically thin and thick cases.

$\log N(C^{\text{iv}})/cm^{-2} = 13.53 \pm 0.07$. No $Si^{\text{iv}}$ absorption is detected for the strongest Ly component, for which we compute a 2σ upper limit of $\log N(Si^{\text{iv}}) < 12.7$.

4.3 The metallicity and ionization properties of the gas

While our $N(H^\text{i})$ measurements of the two gas absorption complexes are subject to large uncertainties because of the saturated Ly component, the range of allowed $N(H^\text{i})$ for each system is constrained to be $\log N(H^\text{i})/cm^{-2} \sim 16–18$. Constraining $H^\text{i}$ column density enables us to subsequently constrain the chemical enrichment level of the gas, based on the observed measurements or limits on the column density of ionic metals. Specifically, the mean elemental abundance of the gas, averaged across all components, is related to the column density ratio of the metal ions $C^{+3}$ and $Si^{+3}$ to $H^0$, following an ionization fraction correction, $\log (C/H) = \log N(C^{\text{iv}})/N(H^\text{i}) - \log (f_{C^{+3}}/f_{H^0})$, and $\log (Si/H) = \log N(Si^{\text{iv}})/N(H^\text{i}) - \log (f_{Si^{+3}}/f_{H^0})$, where $f_{H^0}$, $f_{C^{+3}}$, and $f_{Si^{+3}}$ are the ionization fractions of $H^0$, $C^{+3}$, and $Si^{+3}$ ions, respectively.

For a photoionized gas, the ionization fractions of different species can be predicted given constraints on the ionization parameter of the gas, $U$, which is defined as the number of incident ionizing photon per hydrogen atom. To estimate the necessary ionization fraction corrections, we perform a series of photoionization calculations using CLOUDY v.13.03 (Ferland et al. 2013). We consider a plane parallel slab of gas which is irradiated with the updated Haardt & Madau (2001) ionizing background radiation field (HM05 in CLOUDY) at $z = 2.7$. We assume a solar chemical abundance pattern for the gas, and that the gas is at ionization equilibrium. CLOUDY then calculates the expected ionization fractions and column densities for $H^0$, $C^{+3}$, and $Si^{+3}$ ions over a wide range of gas densities and metallicities. Because the radiation field is fixed, changing the gas density leads to a change in the $U$ parameter.

To estimate the metallicity of the gas, we perform two sets of CLOUDY calculations for each absorber: one for an optically thin gas ($\log N(H^\text{i})/cm^{-2} < 17.2$) and another for
a gas that is optically thick (log \(N(H) / cm^{-2} > 17.2\)) to ionizing photons. The adopted \(N(H)\) for these two sets of calculations are matched to the allowed minimum and maximum \(N(H)\) value of each absorber (see § 4.2). We present an example of cloudy predictions in Figure 5, where in the top panels we plot the expected N(C IV)/N(Si IV) ratio plotted versus \(U\) for a range of gas metallicities (from 0.001 to 0.1 solar metallicity). In this example, the calculation are performed for an optically thin gas with log \(N(H) / cm^{-2} = 16.5\) and optically thick gas with log \(N(H) / cm^{-2} = 17.7\), bracketing the range of allowed \(N(H)\) for Complex A.

It is clear from the top two panels of Figure 5 that the predicted column density ratio of C IV and Si IV are insensitive to gas metallicity over the wide range of \(U\) and metallicities probed, for optically thin and thick gases alike. Therefore, the ionization parameter of the gas can be constrained if the N(C IV)/N(Si IV) ratio is known. Constraining \(U\) then enables us to constrain the expected ionization fractions of C +3 and Si +3 relative to H^0, as shown in the bottom panels of Figure 5.

For absorption complex A near LAEs A1 and A2, our Voigt profile fitting analysis yields a column density ratio of log \(N(C(IV))/N(Si(IV)) = 1.37 \pm 0.11\). Comparing the measured N(C IV)/N(Si IV) to cloudy predictions, we find that the estimated \(U\) parameter of the gas ranges from as low as \(log(U) = -1.8 \pm 0.1\) for the minimum allowed \(N(H)\), log \(N(H) / cm^{-2} = 16.5\), to possibly as high as \(log(U) = -1.2 \pm 0.1\) for the maximum allowed H I column density of log \(N(H) / cm^{-2} = 17.0\). Adopting the range of allowed \(U\) allows us to determine the relative ionization fractions of C +3 and Si +3 to H^0. Adopting the solar chemical abundance pattern from Asplund et al. (2009), we estimate that the metallicity of the gas ranges from as low as [M/H] = −3.7±0.2 for log \(N(H) / cm^{-2} = 17.0\), to no higher than [M/H] = −1.8 ± 0.2 for log \(N(H) / cm^{-2} = 16.5\).

For absorption complex B near LAE B1, the non-detection of Si IV absorption implies a lower limit of log \(N(C(IV))/N(Si(IV)) > 0.83 \pm 0.07\). Based on our cloudy calculations, the observed column density ratio requires \(log(U) > -2.1\), which in turn constrains the gas metallicity to [M/H] < −1.9 for the allowed H I column density range of log \(N(H) / cm^{-2} = 16.5 - 17.8\).

It is instructive to compare the properties of these absorbers to similar H I absorption systems observed at \(z = 2 \pm 3\). For both absorption complexes A and B, the ionization parameter of the gas is constrained to log \(U \geq 2\). This lower limit on \(U\) is consistent with the observed ionization states of Lyman limit systems (LLSs; log \(N(H) / cm^{-2} > 17.2\)) at \(z = 2 \pm 3\), for which \(log(U \geq -3)\) is common (e.g., Fumagalli et al. 2016a; Lehner et al. 2016). Furthermore, the allowed metallicities of the gas, which range from as low as \(\sim 0.001\) solar if the absorbers are optically thick, to no higher than \(\sim 0.02\) solar if the absorbers are in the optically thin regime, are similar to the range of metallicities seen in \(z = 2 \pm 3\) LLSs (e.g., Fumagalli et al. 2016a; Lehner et al. 2016).

5 DISCUSSION

In the previous section, we characterize strong H I absorption systems at \(d < 300\) kpc and \(|\Delta z| < 500\) km s^{-1} from the three faint LAEs. It is also interesting to explore whether the gas is likely to be physically associated with these faint galaxies, and whether the galaxies themselves could be responsible for the origin of the gas. Because of the low completeness of the R08 spectroscopic observations of the field around DMS 2139–0405, it is challenging to link the absorption to a particular galaxy unambiguously, considering that it is possible other faint galaxies occur at smaller \(d\) from the QSO than the nearest (in projection) LAE identified in our study (LAE A1 for Complex A and LAE B1 for Complex B). For the sake of argument, however, we begin our discussion under the assumption that the absorption is associated with the CGM of the LAE that is closest to it in projected distance. We will put this assumption under a closer examination in the next subsection.

5.1 Are the observed metals due to local outflows?

We first investigate whether the gas can originate in a starburst-driven wind originating in the LAEs. Here we focus our discussion on Complex A, which is found closest to LAE A1 at \(d = 87\) kpc. For a wind traveling with a velocity of \(v_{out} = 200\) km s^{-1}, which is typical of \(z \sim 2.5\) LAEs (e.g., Trainor et al. 2015), a minimum of \(\sim 0.4\) Gyr is required for outflowing material to reach a distance of \(\sim 90\) kpc. This simple timing argument rests on the assumption that the outflow velocity can be maintained over cosmic time. In principle, if the galaxy had experienced a starburst episode at around \(z \sim 3.3\) or earlier, enough time would have elapsed for the outflowing material to reach the QSO sightline at the epoch of the observations.

However, energetically it is a non trivial problem to propel gas to large distances from galaxies at high velocity (e.g., Gauthier & Chen 2012). Moreover, while down-the-barrel spectroscopic observations of \(z > 2\) Lyman Break Galaxies (LBGs) and LAEs at have established the ubiquity of galactic-scale winds at high redshifts (e.g., Steidel et al. 2010; Hashimoto et al. 2013; Shibuya et al. 2014; Trainor et al. 2015), a primary caveat of the down-the-barrel method is that the galacto-centric distance of the gas is not well constrained. In the nearby universe, galactic-scale outflows have yet to be directly detected at \(d \gtrsim 10\) kpc from starburst galaxies.

In spite of the lack of direct observation of outflowing gas far from galaxies, kinematic and chemical arguments have been put forward to link outflows with absorbing gas at \(d \gtrsim 50\) kpc from both low- and high-redshift galaxies (e.g., Borthakur et al. 2013; Crighton et al. 2015; Muzahid et al. 2015). In particular, outflowing gas is expected to exist, the implication is that the large majority of the H I-bearing gas would be even more metal-poor than inferred here, and the conclusion that the observed gas complex is predominantly metal-poor (\(\lesssim 0.01\) solar) would remain valid. In that regard, the metallicity constraints obtained in this section should be viewed as conservative upper limits to the mean chemical enrichment level in the bulk of the gas.

\[3\] For our metallicity estimates, all the C IV and Si IV absorption are attributed to the total observed H I absorption. It is possible that these ionic metals originate in only a small but metal-rich portion of the H I-bearing gas complex. If such metal-rich pockets
exhibit a large velocity width and a high chemical enrichment level, due to the supernova origin of a super-galactic wind. Observations of blue-shifted ISM absorption lines in lensed LBGs at \(z \sim 2\) demonstrate that a high metallicity of 0.4 – 0.8 solar is common in outflows (e.g., Pettini et al. 2002; Dessauges-Zavadsky et al. 2010).

While no direct chemical abundance measurement is available for LAE A1, the lack of continuum counterpart of this LAE (\(r > 26.6, K_s > 24.5\)) implies that the host galaxy is intrinsically faint and low-mass (\(L \lesssim 0.1 L^\odot, M_{\text{star}} \lesssim 10^8 M^\odot\)). Based on the mass-metallicity relation at \(z \sim 2\), the expected gas-phase metallicity in the ISM of a \(M_* \sim 10^7 M^\odot\) galaxy is 0.1 – 0.2 solar (e.g., Steidel et al. 2014; Sanders et al. 2015). The expected ISM abundance is similar to the mean metallicity of 0.2 ± 0.1 solar estimated from nebular emission lines in stacked spectra of faint \(z \sim 2.5\) LAEs (Trainor et al. 2016).

In contrast, the inferred metallicity of Complex A is < 0.02 solar, which is at least an order of magnitude lower than both the observed metallicities in high-redshift outflows and the inferred chemical abundance in the ISM of faint LAEs. The stark difference between the chemical enrichment level of the gas and those expected in outflowing material and ISM gas indicates that a starburst-driven wind is unlikely to be the origin of the observed gas.

Next, we examine whether the observed gas in Complex A is likely to be situated in the CGM of LAE A1. The total carbon mass enclosed within a projected distance \(d\) in the CGM is \(M_C = \pi d^2 \langle N(C^{\text{IV}}) \rangle 12 m_{\text{H}} f_{\text{cov}} / f_{C^{\text{IV}+3}}\), where \(\langle N(C^{\text{IV}}) \rangle\) is the mean \(C^{\text{IV}}\) column density, \(f_{\text{cov}}\) is the mean gas covering fraction and \(f_{C^{\text{IV}+3}}\) is the ionization fraction of triply ionized carbon. Assuming that that the radial profile of \(C^{\text{IV}}\) absorption within \(d \sim 90\) kpc has a mean column of \(\langle N(C^{\text{IV}}) \rangle = 10^{14.5} \text{cm}^{-2}\) (which is the measured total \(C^{\text{IV}}\) column density in Complex A), and adopting \(f_{C^{\text{IV}+3}} = 0.2\) based on our CLOUDY calculations, we estimate a total CGM carbon mass of \(M_C \sim 2 \times 10^8 M^\odot (f_{\text{cov}})\). In a similar fashion, the observed \(N(S^{\text{II}V})\) implies a total silicon mass of \(M_{Si} \sim 3 \times 10^8 M^\odot (f_{\text{cov}})\). For \(f_{\text{cov}} = 0.5\), which is typical at similar \(d\) from both low-redshift dwarf/sub-L^* galaxies and high-redshift LBGs (e.g., Bordoloi et al. 2014; Liang & Chen 2014; Rudie et al. 2019), the estimated total carbon and silicon masses in the CGM are \(\sim 1 \times 10^8 M^\odot\) and \(\sim 1.5 \times 10^8 M^\odot\), respectively.

Can the galaxy produce the estimated amount of CGM metals? Assuming the galaxy has maintained a constant star-formation rate of 0.1 \(M^\odot\) yr\(^{-1}\) (see §3) over time, and adopting a supernova carbon yield of 0.0088 \(M^\odot\) per solar mass of star formation (Peeples et al. 2014), we estimate that a total mass of \(\sim 9 \times 10^4 M^\odot\) in carbon is produced over a \(\sim 0.1\) Gyr timescale. The recent metal production budget in the galaxy is far below the estimated total carbon mass to explain the observed metal enrichment of the surrounding IGM through local galactic feedback. This exercise suggests that the absorbing gas is unlikely to originate in any form of CGM around the individual LAEs.

5.2 Evidence for accretion from a cosmic web filament?

In Figure 6, we present gas-phase metallicity measurements versus \(d\) for \(z > 2\) galaxies, including LAEs in our study. For comparison, we indicate the range of metallicities seen in the \(z \sim 2.5\) IGM in the shaded gray area (Simcoe et al. 2004). The inferred low metallicities of absorbers in our study (no higher than 0.02 solar and possibly lower than 0.001 solar metallicity) indicate that the gas originates in the more metal-poor IGM. Indeed, chemical enrichment level in the IGM at \(z \sim 2.5\) is characterized by a median metallicity of [M/H] = −2.8 (or 0.002 solar) and a dispersion of 0.75 dex (e.g., Schaye et al. 2003; Simcoe et al. 2004), in excellent agreement with the metallicity constraints for absorbers in our study.

Furthermore, the presence of multiple LAEs within a close proximity from each other along the slit implies that the emitters may be situated in a coherent structure. Using the field LAE luminosity function at \(z \sim 3\) from Drake et al. (2017), we expect to detect only \(\sim 0.3\) LAEs above the detection limit in the observations of R08, within a cosmological volume of \(\sim 30\) Mpc\(^3\) defined by a slit width of 2 arcsec, a slit length of 453 arcsec, and a redshift range of \(z = 2.74 – 2.78 (= 3200\) km s\(^{-1}\) in velocity space, bracketing the redshifts of the Complex A and Complex B). Within this volume, the probability that three LAEs occur by ran-
dom chance is very low, $\sim 3 \times 10^{-3}$. The probability that the grouping of three LAEs should occur within projected $\pm 300$ kpc from each other is even lower, $\sim 1 \times 10^{-4}$, thereby suggesting that the LAEs are part of a spatially coherent structure.

The most plausible explanation for the unique arrangement of Lyα emitters and absorbers in our study (Figure 1) is that the spectrograph slit in R08 was serendipitously aligned (at least partially) with a cosmic web filament at $z \sim 2.8$, thereby explaining the presence of multiple LAEs and Lyα absorbers along a narrow line extending more than 200 kpc in projection (see also Fumagalli et al. 2016b). Several features of the absorbing gas in Complex A near LAEs A1 and A2 provide further support for this scenario.

First, the observed line-of-sight velocity spread of the multi-component CIV absorption in Complex A is low, with a statistical dispersion of $\sigma = 32$ km s$^{-1}$ (see bottom panel of Figure 3). This velocity width is significantly smaller than those observed in down-the-barrel outflows from $z = 2 - 3$ galaxies ($\gtrsim 100$ km s$^{-1}$, e.g., Trinquier et al. 2015; Jones et al. 2018) and expected from dynamical motion in the CGM (the line-of-sight velocity dispersion for virialized motion is $\sigma_b \approx 60 - 130$ km s$^{-1}$ in a $M_h = 10^{11-12} M_\odot$ halo). However, the observed narrow CIV width is consistent with the expected velocity spread for a quiescent, $\sim 100$ kpc scale gaseous filament that is simply moving with the Hubble flow at this redshift. Second, the absorption is flanked on opposite sides of the slit by the two LAEs (top panel of Figure 3), which are separated by a projected 248 kpc and are consistent within uncertainties with having zero systemic velocity offset from each other. Furthermore, the mean Doppler parameter for the CIV absorption components in Complex A is $\langle b \rangle = 9.8 \pm 1.0$ km s$^{-1}$, which is very similar to the mode of the CIV Doppler parameter distribution predicted by simulations of gas accretion in cosmic filaments (e.g., Rauch et al. 1997). Together, these kinematic and morphological features suggest that the absorbing gas originates in a large and dynamically cold structure, such as that of a gaseous filament.

Recent hydrodynamical simulations of galaxy formation predict that cosmological accretion from the IGM onto high-redshift galaxies proceed along filaments that often manifest as low-metallicity LLSs (e.g., Fumagalli et al. 2011; van de Voort et al. 2012; Faucher-Giguère et al. 2015), in agreement with the observed properties of the absorption complexes in our study. In this physical picture, infalling streams of gas can proceed along such filamentary structure and fuel star-formation activity in the LAEs. This feeding scenario is further supported by the prominent blue peak that is observed in the Lyα emission line profile of LAE A1 (R08), an indication that the Lyman-emitting gas is flowing onto the galaxy (e.g., Verhamme et al. 2006). In conclusion, the observed properties of the absorbers and the unique orientation of faint LAEs and the background sightline are in good agreement with the predictions of infalling IGM gas along cosmic filaments.

### 6 SUMMARY

Using high-resolution optical echelle spectrum of QSO DMS 2139–0405 we studied the gaseous environment of three faint LAEs ($L_{Ly\alpha} \approx 2 \times 10^{41}$ erg s$^{-1}$) at $z = 2.7 - 2.8$. Strong H I absorption systems are found at $d < 300$ kpc and $|\Delta v| < 500$ km s$^{-1}$ from the LAEs, allowing us to investigate the connection between faint, low-mass galaxies ($L < 0.1 L^*$, $M_{star} < 10^8 M_\odot$) and the IGM at this redshift. Our findings are summarized below.

(i) We find that the absorbers have H I column densities which are consistent with being partial or full LLSs, $\log N$(H I)/cm$^{-2} = 16 - 18$. The observed metal abundances of the gas are low, no higher than 0.02 solar metallicity if the absorbers are optically thin and possibly as low or lower than $\sim 0.001$ solar metallicity if the absorbers are optically thick LLSs. The inferred metallicities and ionization states of the gas are comparable to what have been observed in $z = 2 - 3$ LLSs.

(ii) Focusing on Complex A, which is found at $d = 87$ and 161 kpc from LAEs A1 and A2, we examined the possibility that the gas originates in a locally enriched CGM around one of the galaxies. While a galactic wind with a constant velocity of 200 km s$^{-1}$ may be formally able to cover a distance of $\sim 90$ kpc and reach the QSO sightline in $\sim 0.4$ Gyr, direct evidence for the existence of outflows at large distances remains elusive (see § 1). Furthermore, given the inferred low SFR of the LAE, the expected metal production budget is insufficient to explain the observed metals as having been produced and expelled by these galaxies within the last $\sim 0.1$ Gyr.

(iii) The low metallicities of the gas are consistent with the gas originating in the general IGM. The spatial alignment of the LAEs along the slit, and the alignment of said LAEs and the absorbers in velocity space, suggest that we may be looking perpendicularly at a filamentary structure spanned by the galaxies over at least $\sim 250$ kpc in projection. The range of $N$(H I) of the gas, its low metallicity and line-of-sight velocity dispersion, together with the observed mean CIV Doppler parameter values and presence of faint galaxies in the vicinity of the absorbers, are all in good agreement with the predictions for infalling gas along cosmic filaments (Rauch, Haehnelt & Steinmetz 1997; Fumagalli et al. 2011, van de Voort et al. 2012). The prominence of a blue Lyα emission peak in one of the LAEs (R08) is consistent with radiative transfer through infalling gas, making this a rare occasion where we can observe an accretion stream together with its final destination.

We expect that future spectroscopic surveys of QSO fields in search of faint line emitters, taking advantage of wide-field integral field spectrographs and coupled with very deep imaging of these fields, will provide new insights into the gaseous environment of low-mass galaxies at high redshifts and offer critical information on the relative importance and roles of feeding and feedback during the cosmic high noon.

### ACKNOWLEDGMENTS

FSZ acknowledges generous support from the Brinson Foundation and the Observatories of the Carnegie Institu-
REFERENCES

Adelberger, K. L., Steidel, C. C., Shapley, A. E., & Pettini, M. 2003, ApJ, 584, 45
Asplund, M., Grevesse, N., Sauval, A. J., & Scott, P. 2009, ARA&A, 47, 481
Becker, G. D., Sargent, W. L. W., Rauch, M., & Calverley, A. P. 2011, ApJ, 735, 93
Bernstein, R., Shectman, S. A., Gunnels, S. M., Mochnacki, S., & Athey, A. E. 2003, Proc. SPIE, 4841, 1694
Biellby, R. M., Tummuangpak, P., Shanks, T., et al. 2016, MNRAS, 456, 4061
Bordoloi, R., Tumlinson, J., Werk, J. K., et al. 2014, ApJ, 796, 136
Borthakur, S., Heckman, T., Strickland, D., Wild, V., & Tumlinson, J. 2014, ApJ, 799, 138
Brocklehurst, M. 1971, MNRAS, 153, 471
Carswell, R. F., & Webb, J. K. 2011, MNRAS, 418, 1115
Cassata, P., Le Fèvre, O., Garilli, B., et al. 2011, A&A, 525, A143
Chen, S.-F. S., Simcoe, R. A., Torrey, P., et al. 2017, ApJ, 850, 188
Cooper, T. J., Simcoe, R. A., Cooksey, K. L., et al. 2019, arXiv:1901.05980
Crighton, N. H. M., Hennawi, J. F., & Prochaska, J. X. 2013, ApJ, 776, L18
Crighton, N. H. M., Hennawi, J. F., Simcoe, R. A., et al. 2015, MNRAS, 446, 18
Dessauges-Zavadsky, M., D’Odorico, S., Schaerer, D., et al. 2010, A&A, 510, A26
Drake, A. B., Guiderdoni, B., Blaizot, J., et al. 2017, MNRAS, 471, 267
Erb, D. K., Steidel, C. C., Trainor, R. F., et al. 2014, ApJ, 795, 33
Faucher-Giguère, C.-A., & Kereš, D. 2011, MNRAS, 412, L118
Faucher-Giguère, C.-A., Hopkins, P. F., Kereš, D., et al. 2015, MNRAS, 449, 987
Fumagalli, M., Prochaska, J. X., Kasen, D., et al. 2011, MNRAS, 418, 1796
Fumagalli, M., O’Meara, J. M., & Prochaska, J. X. 2016, MNRAS, 455, 4100
Fumagalli, M., Cantalupo, S., Dekel, A., et al. 2016, MNRAS, 462, 1978
Ferland, G. J., Porter, R. L., van Hoof, P. A. M., et al. 2013, Rev. Mexicana Astron. Astrofís., 49, 137
Gauthier, J.-R., & Chen, H.-W. 2012, MNRAS, 424, 1952
Goerdt, T., Dekel, A., Sternberg, A., et al. 2010, MNRAS, 407, 613
Haardt, F., & Madau, P. 2001, Clusters of Galaxies and the High Redshift Universe Observed in X-rays, Haehnelt, M. G., Steinmetz, M., & Rauch, M. 1995, ApJ, 465, L95
Hafen, Z., Faucher-Giguère, C.-A., Anglés-Alcázar, D., et al. 2017, MNRAS, 469, 2292
Hashimoto, T., Ouchi, M., Shimasaku, K., et al. 2013, ApJ, 765, 70
Hashimoto, T., Verhamme, A., Ouchi, M., et al. 2015, ApJ, 812, 157
Jones, T., Stark, D. P., & Ellis, R. S. 2018, ApJ, 863, 191
Kawata, D., & Rauch, M. 2007, ApJ, 663, 38
Kennicutt, R. C., & Evans, N. J. 2012, ARA&A, 50, 531
Kereš, D., Katz, N., Weinberg, D. H., & Davé, R. 2005, MNRAS, 363, 2
Konno, A., Ouchi, M., Nakajima, K., et al. 2016, ApJ, 823, 20
Lehner, N., O’Meara, J. M., Howk, J. C., Prochaska, J. X., & Fumagalli, M. 2016, ApJ, 833, 283
Liang, C. J., & Chen, H.-W. 2014, MNRAS, 445, 2061
Marchesini, D., van Dokkum, P. G., Förster Schreiber, N. M., et al. 2009, ApJ, 701, 1765
Muzahid, S., Kacprzak, G. G., Churchill, C. W., et al. 2015, ApJ, 811, 132
Oke, J. B., Cohen, J. G., Carr, M., et al. 1995, PASP, 107, 375
Peeples, M. S., Werk, J. K., Tumlinson, J., et al. 2014, ApJ, 786, 54
Persson, S. E., Barkhouse, R., Birk, C., et al. 2008, Proc. SPIE, 7014, 70142V
Pettini, M., Rix, S. A., Steidel, C. C., et al. 2002, ApJ, 569, 742
Porciani, C., & Madau, P. 2005, ApJ, 625, L43
Rauch, M., Haehnelt, M. G., & Steinmetz, M. 1997, ApJ, 481, 601
Rauch, M., Sargent, W. L. W., & Barlow, T. A. 2001, ApJ, 554, 823
Rauch, M., Sargent, W. L. W., Barlow, T. A., & Carswell, R. F. 2001, ApJ, 562, 76
Rauch, M., Haehnelt, M., Bunker, A., et al. 2008, ApJ, 681, 856-880
Rauch, M., Becker, G. D., Haehnelt, M. G., et al. 2011, MNRAS, 418, 1115
Rauch, M., Becker, G. D., & Haehnelt, M. G. 2016, MNRAS, 455, 3991
Reddy, N. A., Steidel, C. C., Pettini, M., et al. 2008, ApJS, 175, 48
Reddy, N. A., Pettini, M., Steidel, C. C., et al. 2012, ApJ, 754, 25
Rosdahl, J., & Blaizot, J. 2012, MNRAS, 423, 344
Rudie, G. C., Steidel, C. C., Trainor, R. F., et al. 2012, ApJ, 750, 67
Rudie, G. C., Steidel, C. C., Pettini, M., et al. 2019, arXiv:1903.00004
Ryan-Weber, E. V., Pettini, M., Madau, P., & Zych, B. J. 2009, MNRAS, 395, 1476
Sanders, R. L., Shapley, A. E., Kriek, M., et al. 2015, ApJ, 799, 138
Scannapieco, C., Tissera, P. B., White, S. D. M., & Springel, V. 2005, MNRAS, 364, 552

MNRAS 000, 1–13 (2019)
IGM gas accretion onto $z \sim 2.8$ Ly$\alpha$ emitters

Schaye, J., Aguirre, A., Kim, T.-S., et al. 2003, ApJ, 596, 768
Shen, S., Madau, P., Guedes, J., et al. 2013, ApJ, 765, 89
Shibuya, T., Ouchi, M., Nakajima, K., et al. 2014, ApJ, 788, 74
Simcoe, R. A., Sargent, W. L. W., & Rauch, M. 2004, ApJ, 606, 92
Simcoe, R. A., Sargent, W. L. W., Rauch, M., & Becker, G. 2006, ApJ, 637, 648
Simcoe, R. A., Cooksey, K. L., Matejek, M., et al. 2011, ApJ, 743, 21
Skrutskie, M. F., Cutri, R. M., Stiening, R., et al. 2006, AJ, 131, 1163
Songaila, A. 2001, ApJ, 561, L153
Stalder, B., Stark, A. A., Amato, S. M., et al. 2014, Proc. SPIE, 9147, 91473Y
Steidel, C. C., Erb, D. K., Shapley, A. E., et al. 2010, ApJ, 717, 289
Steidel, C. C., Rudie, G. C., Strom, A. L., et al. 2014, ApJ, 795, 165
Strom, A. L., Steidel, C. C., Rudie, G. C., Trainor, R. F., & Pettini, M. 2018, ApJ, 868, 117
Trainor, R. F., Steidel, C. C., Strom, A. L., & Rudie, G. C. 2015, ApJ, 809, 89
Trainor, R. F., Strom, A. L., Steidel, C. C., & Rudie, G. C. 2016, ApJ, 832, 171
van de Voort, F., Schaye, J., Altay, G., & Theuns, T. 2012, MNRAS, 421, 2809
Verhamme, A., Schaerer, D., & Maselli, A. 2006, A&A, 460, 397
Wiersma, R. P. C., Schaye, J., Dalla Vecchia, C., et al. 2010, MNRAS, 409, 132
Wild, V., Kauffmann, G., White, S., et al. 2008, MNRAS, 388, 227
York, D. G., Adelman, J., Anderson, J. E., Jr., et al. 2000, AJ, 120, 1579

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APPENDIX A: OTHER ABSORPTION SYSTEMS AT $Z = 2.7 - 2.8$.

Two additional Ly$\alpha$ absorption complexes are found at $z = 2.7 - 2.8$ along the spectrum of background QSO DMS 2139−0405 (see Figure 1): Complex X at $z = 2.7147$ and Complex Y at $z = 2.7863$. No known LAE is found within $d < 300$ kpc and $|\Delta v| < 500$ km s$^{-1}$ from these Ly$\alpha$ absorption systems.

A strong, saturated Ly$\alpha$ component dominates the total H$\textsc{i}$ column density of Complex X (Figure A1), with a best-fit log $N$(H$\textsc{i}$)/cm$^{-2} = 17.9^{+0.2}_{-0.3}$. For this saturated system, we estimate that the range of allowed $N$(H$\textsc{i}$)/cm$^{-2}$ is $15.7 - 18.2$ at 95 percent confidence level. In addition to Ly$\alpha$, C$\textsc{iv}$ absorption is observed in Complex X as well, with a measured total column density of log $N$(C$\textsc{iv}$) = $13.8 \pm 0.11$. Unfortunately, it is not possible to constrain the Si$\textsc{iv}$ column density, due to contamination from Ly$\alpha$ forest lines at higher redshifts.

Only Ly$\alpha$ absorption is detected in Complex Y (Figure A2). The range of allowed $N$(H$\textsc{i}$) for this saturated system is log $N$(H$\textsc{i}$)/cm$^{-2} = 16.7 - 17.8$ at 95% confidence level. For the absence of ionic metal absorption, we calculate upper limits on log $N$(C$\textsc{iv}$)/cm$^{-2} < 13.2$ and log $N$(Si$\textsc{iv}$)/cm$^{-2} < 12.6$ for C$\textsc{iv}$ and Si$\textsc{iv}$, respectively.

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Figure A2. Continuum-normalized absorption profiles of Complex Y at $z = 2.7863$. The absorption spectra and 1-$\sigma$ error array are shown in black and cyan histograms, whereas the best-fitting Voigt profile models are shown in red curves. The magenta tick marks indicate the location of individual components. Contaminating features are dotted out for clarity.

Table A1. Voigt profile fitting results for additional absorption complexes at $z = 2.7 - 2.8$

| Species | $z$ | $\log N_c$/cm$^{-2}$ | $b_c$ (km s$^{-1}$) |
|---------|----|----------------------|---------------------|
| Complex X at $z = 2.7147$ | | | |
| $\text{H}^\text{i}$ | 2.7122 ± 0.0001 | 13.82 ± 0.12 | 57.0 ± 10.7 |
| $\text{CIV}$ | ... | < 13.0 | ... |
| $\text{SiIV}$ | ... | < 12.7 | ... |
| $\text{H}^\text{i}$ | 2.71466 ± 0.00003 | 17.9$^{+0.2}_{-0.3}$ | 47$^{+1}_{-1}$ |
| $\text{CIV}$ | 2.71469 ± 0.00003 | 13.58 ± 0.14 | 13.7 ± 3.9 |
| $\text{SiIV}$ | 2.71519 ± 0.00015 | 13.50 ± 0.18 | 31.8 ± 13.5 |
| $\text{H}^\text{i}$ | 2.71710 ± 0.00037 | 13.73 ± 0.25 | 94.3 ± 44.3 |
| $\text{CIV}$ | ... | < 13.1 | ... |
| $\text{H}^\text{i}$ | 2.71904 ± 0.00005 | 13.59 ± 0.09 | 35.0 ± 8.0 |
| $\text{CIV}$ | ... | < 12.9 | ... |
| $\text{H}^\text{i}$ | 2.71974 ± 0.00001 | 13.71 ± 0.13 | 11.2 ± 2.0 |
| $\text{CIV}$ | ... | < 12.7 | ... |
| $\text{H}^\text{i}$ | 2.72037 ± 0.00002 | 13.49 ± 0.12 | 8.3 ± 2.0 |
| $\text{CIV}$ | ... | < 12.9 | ... |
| $\text{H}^\text{i}$ | 2.72086 ± 0.00006 | 13.15 ± 0.11 | 23.0 ± 7.1 |
| $\text{CIV}$ | ... | < 13.1 | ... |
| $\text{H}^\text{i}$ | 2.72217 ± 0.00003 | 12.99 ± 0.10 | 12.3 ± 3.8 |
| $\text{CIV}$ | ... | < 13.0 | ... |
| $\text{SiIV}$ | ... | < 12.4 | ... |

Complex Y at $z = 2.7863$

| $\text{H}^\text{i}$ | 2.78630 ± 0.00002 | 17.6$^{+0.1}_{-0.6}$ | 27$^{+3}_{-1}$ |
| $\text{CIV}$ | ... | < 12.9 | ... |
| $\text{SiIV}$ | ... | < 12.3 | ... |
| $\text{H}^\text{i}$ | 2.78790 ± 0.00002 | 14.34 ± 0.15 | 24.0 ± 3.2 |
| $\text{CIV}$ | ... | < 12.8 | ... |
| $\text{SiIV}$ | ... | < 12.2 | ... |