THE PHYSICAL PROPERTIES OF HIGH-MASS STAR-FORMING CLUMPS: A SYSTEMATIC COMPARISON OF MOLECULAR TRACERS

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ABSTRACT

We present observations of HCO+ and H13CO+, N2H+, HCS+, HNC and HN13C, SO and 34SO, CCH, SO2, and CH3OH-E toward a sample of 27 high-mass clumps coincident with water maser emission. All transitions are observed with or convolved to nearly identical resolution (30″), allowing for inter-comparison of the clump properties derived from the mapped transitions. We find that N2H+ emission is spatially differentiated compared with the dust and the other molecules toward a few very luminous cores (10 of 27) and the N2H+ integrated intensity does not correlate well with dust continuum flux. We calculate the effective excitation density, \( n_{\text{eff}} \), the density required to excite a 1 K line in \( T_{\text{kin}} = 20 \) K gas for each molecular tracer. The intensity of molecular tracers with larger effective excitation densities (\( n_{\text{eff}} \gtrsim 10^5 \text{ cm}^{-3} \)) appears to correlate more strongly with the submillimeter dust continuum intensity. The median sizes of the clumps are anti-correlated with the \( n_{\text{eff}} \) of the tracers (which span more than three orders of magnitude). Virial mass is not correlated with \( n_{\text{eff}} \), especially where the lines are optically thick as the linewidths may be broadened significantly by non-virial motions. The median mass surface density and median volume density of the clumps are correlated with \( n_{\text{eff}} \) indicating the importance of understanding the excitation conditions of the molecular tracer when deriving the average properties of an ensemble of cores.

Key words: dust, extinction – ISM: clouds – stars: formation

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1. INTRODUCTION

Throughout their short lives, massive stars make significant contributions to their host galaxy’s energy budget. From strong winds and intense radiation during their brief main-sequence existence to their spectacular end in powerful supernovae, massive stars provide substantial energetic feedback (Motte & Hennebelle 2008). On galactic scales, massive stars enrich the interstellar medium (ISM) with heavier elements, and their powerful winds may trigger successive episodes of star formation. Understanding the role of massive stars in galaxy evolution is essential for improving galaxy evolution models and properly treating feedback from high-mass stellar evolution. Despite their importance, many fundamental questions remain to be answered about the formation of massive stars.

While the theory of low-mass star formation is relatively mature (e.g., Shu et al. 1987), it does not simply extend to the formation of stars above \( \sim 8 \ M_\odot \); indeed, the basic formation mechanisms for high-mass stars are still debated. Three competing theories attempt to describe the origin of massive stars (Zinnecker & Yorke 2007). The core accretion model essentially scales the theory of low-mass star formation to higher masses by assuming that massive stars form from massive cores (Zinnecker & Yorke 2007). Core accretion models, such as those studied by Krumholz et al. (2009), require non-spherical accretion to overcome the intense radiation pressure that massive protostars create (and that causes theories for low-mass star formation to fail for larger masses). Core accretion models must explain why a massive cloud would collapse to form a single, massive core, rather than fragment into several less-massive cores. Recent work by Krumholz (2006) and Krumholz et al. (2007a, 2007b) suggests that radiation feedback may solve this problem by changing the equation of state of the gas, which determines whether the clump will fragment, and thus allows a more massive core to collapse.

Competitive accretion models, on the other hand, assume that fragmentation produces many low-mass stellar cores while high-mass stars form subsequently due to continued accretion (e.g., Bonnell & Bate 2006). This model may also explain why massive stars are found preferentially in the center of clusters—residing deep in the cluster potential allows cores to grow more massive, providing greater gravity to attract the gas, thus leading to runaway accretion (Bonnell et al. 2001). It is unclear whether this mass segregation is a result of how massive stars form (Bonnell & Davies 1998; Huff & Stahler 2006) or if it is a consequence of cluster dynamics (Tan et al. 2006; McMillan et al. 2007). For the densest star clusters, massive stars may form via stellar collisions, although this process requires densely packed objects of relatively high mass to proceed (Zinnecker & Yorke 2007).

Current observational constraints are not sufficiently quantitative to discriminate between models of high-mass star formation (Krumholz & Bonnell 2007; Motte & Hennebelle 2008). Theoretical models are fundamentally limited by an incomplete understanding of the properties of the dense gas in high-mass (\( M > \) few \( M_\odot \)) star-forming clumps. While the high spatial resolution available with ALMA is required to observe massive star formation on the scale of individual stars (Krumholz & Bonnell 2007), on the scale of individual star clusters, progress can be made by improving our understanding of the properties.
of massive clumps. Recent mapping surveys of dust continuum emission (Mueller et al. 2002; Beuther et al. 2002; Williams et al. 2005) as well as molecular emission from multiple transitions of CS and HCN (Plume et al. 1992, 1997; Shirley et al. 2003; Wu et al. 2010) have attempted to characterize the large-scale density and temperature structure of the massive progenitors of clusters. Blind surveys of continuum emission in the galactic plane (e.g., GLIMPSE, Benjamin et al. 2005; BGPS, Aguirre et al. 2011; ATLASGAL, Schuller et al. 2009; HIGAL, Molinari et al. 2010; etc.) reveal new sites of potential star formation. Molecular line observations are essential for determining the basic properties of these newly discovered sources. From molecular line observations, we can calculate kinematic distances and derive physical properties, such as size, mass, and density, of these sources. Proper interpretation of those results requires that we understand how the derived clump properties depend on the tracer and the excitation conditions required to excite that transition.

In this paper, we analyze observations of a sample of 27 high-mass clumps associated with H2O masers in the dense gas tracers N2H+, HCO+, HCS+, HNC, CCH, SO, SO2, and CH3OH in a total of 12 transitions among the isotopologues. We study spatial information in emission maps of transitions of N2H+, HCO+, HNC, CCH, SO, and for previously published CS transitions. This particular set of transitions covers more than three orders of magnitude in effective excitation density with nearly identical resolution (30") for all twelve transitions. Matching the resolution between tracers is of critical importance for systematic statistical comparisons of the properties of the sources. The transitions studied in this work represent diverse probes of the chemistry and physical structure in these clumps. They include molecular ions, sulfur-bearing species, the simplest carbon chain molecule, and the isomer of a commonly observed dense gas tracer. The rich chemistry in high-mass star-forming regions may provide clues to the evolutionary state of the clump and there is a rich literature of line surveys studying the chemistry in high-mass star-forming regions in detail (e.g., Sakai et al. 2008; Beuther et al. 2009; Vasyunina et al. 2011). Though many of our molecular species have been observed in other studies, we do not study the chemistry of these clouds in detail here. Instead, we use the results of this mapping survey to study in detail how fundamental properties of the clumps (e.g., size, mass, surface density, etc.) vary with the effective excitation density of the tracer.

In our analysis we use the effective density $n_{\text{eff}}$, rather than critical density $n_{\text{crit}}$, to characterize the environment where a transition is excited. As discussed in Evans (1999), the ability to detect a transition is usually taken to mean that $n > n_{\text{crit}}$, but in truth lines are easily excited in subthermally populated gas with densities more than an order of magnitude lower than $n_{\text{crit}}$. Molecular filling fraction, line frequency, optical depth, multilevel excitation effects, and trapping can all affect the observed line strength. To account for these effects, we use the Evans (1999) definition of effective density as the density needed to excite a $T_R = 1$ K line for a given kinetic temperature as calculated with a non-LTE radiative transfer code. In this paper, we update and expand the table of effective densities given in Evans (1999). We use the online version of RADEX (van der Tak et al. 2007), a non-LTE radiative transfer code, assuming $\log(N/\Delta v) = 13.5 \text{ cm}^{-2}/(\text{km s}^{-1})$ and a kinetic temperature of 20 K for all species (see Table 1, Figure 1).

The sample of sources is drawn from the Plume et al. (1992, 1997) survey of multiple CS transitions (see Table 2). Most of these sources have been mapped and modeled in 350 $\mu$m dust continuum emission (Mueller et al. 2002). Maps of CS $J = 5 \rightarrow 4$ emission offer additional constraints on size and virial mass of these dense clumps (Shirley et al. 2003). Wu et al. (2010) extended the mapped data set to include two CS transitions ($J = 7 \rightarrow 6$ and $J = 2 \rightarrow 1$) as well as adding two transitions of HCN ($J = 1 \rightarrow 0$ and $J = 3 \rightarrow 2$). Our sample of 27 sources represents a range of bolometric luminosities from $1.1 \times 10^3 L_\odot$ to $1.6 \times 10^6 L_\odot$ (median $4.4 \times 10^4 L_\odot$). CS clump sizes are $\sim 0.25$ pc on average. The median virial mass, calculated from CS $J = 5 \rightarrow 4$ emission in (Shirley et al. 2003) is $600 M_\odot$. Sources are at distances ranging from 0.7 kpc to 8.5 kpc (median 2.8 kpc) so that with a 30") single dish beam we have a corresponding sensitivity to spatial scales of 0.05–0.62 pc, respectively (median 0.21 pc). All of our maps have nearly identical resolution (30"), allowing for inter-comparison of the source properties derived from those maps. The diversity of this

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### Table 1

Properties of Observed Transitions

| Transition | $v$ (GHz) | $\theta_{\text{ab}}$ (°) | $g_a$ | $E_a/k$ (K) | $A_{\text{rot}}$ (s$^{-1}$) | $n_{\text{eff}}$ (20 K) | $n_{\text{crit}}$ (20 K) |
|------------|-----------|--------------------------|-------|------------|--------------------------|--------------------------|--------------------------|
| N$_2$H$^+$ 1–0 | 93.1737767 | 18 | 3 | 4.47 | 3.63 $\times$ 10$^{-5}$ | 2.0 $\times$ 10$^3$ | 1.6 $\times$ 10$^3$ |
| H$_2$CO$^+$ 3–2 | 260.2553390 | 27.9 | 7 | 24.98 | 1.34 $\times$ 10$^{-3}$ | 3.2 $\times$ 10$^6$ | 3.2 $\times$ 10$^6$ |
| HCO$^+$ 3–2 | 267.5576190 | 27.1 | 7 | 25.68 | 1.48 $\times$ 10$^{-3}$ | 2.3 $\times$ 10$^6$ | 3.5 $\times$ 10$^6$ |
| N$_2$H$^+$ 3–2 | 279.5117010 | 26.0 | 7 | 26.83 | 1.26 $\times$ 10$^{-3}$ | 2.7 $\times$ 10$^4$ | 3.0 $\times$ 10$^4$ |
| HCS$^+$ 6–5 | 256.0271000 | 28.3 | 13 | 43.01 | 3.46 $\times$ 10$^{-4}$ | ... | 9.1 $\times$ 10$^5$ |
| CCH 3–2 | 262.07167 | 28.8 | 7 | 21.32 | 5.75 $\times$ 10$^{-5}$ | 1.5 $\times$ 10$^4b$ | 5.15 $\times$ 10$^4b$ |
| HNC 3–2 | 271.9811420 | 27.7 | 7 | 26.11 | 9.34 $\times$ 10$^{-4}$ | 1.7 $\times$ 10$^5$ | 6.2 $\times$ 10$^5$ |
| HN$^{13}$C 3–2 | 261.2633101 | 28.9 | 7 | 25.08 | 6.48 $\times$ 10$^{-4}$ | ... | 4.3 $\times$ 10$^7$ |
| CS 5–4 | 244.9355565 | 29.6 | 11 | 35.27 | 2.98 $\times$ 10$^{-4}$ | 3.7 $\times$ 10$^5$ | 7.3 $\times$ 10$^5$ |
| $^{34}$SO 5–4c | 241.0161940 | 30.1 | 11 | 34.70 | 2.84 $\times$ 10$^{-4}$ | 6.9 $\times$ 10$^6$ | 6.9 $\times$ 10$^6$ |
| CS 6–5 | 261.84296479 | 28.8 | 11 | 47.551 | 2.24 $\times$ 10$^{-4}$ | 2.4 $\times$ 10$^6$ | 1.2 $\times$ 10$^6$ |
| $^{34}$SO 7–6c | 256.0774563 | 29.4 | 15 | 36.26 | 2.19 $\times$ 10$^{-4}$ | ... | 1.2 $\times$ 10$^6$ |
| SO$_2$ 13.5–12.6 | 257.0999658 | 29.3 | 29 | 47.84 | 1.13 $\times$ 10$^{-4}$ | ... | 3.42 $\times$ 10$^5$ |
| CH$_3$OH-E 21–10 | 261.805736 | 28.8 | 5 | 23.30 | 6.71 $\times$ 10$^{-5}$ | ... | 4.0 $\times$ 10$^6$ |

**Notes.**

*a* Average frequency for the blended hyperfine components.

*b* Assuming H$_3$C$^+$ $J = 3 \rightarrow 2$ collision rates since no data are available for CCH.

*c* These transitions were observed in Shirley et al. 2003, but are presented here for completeness.
| Source   | α (J2000.0)  | δ (J2000.0)  | D  | $L_{bol}$ ($L_\odot$) | $T_{iso}$ (K) | $p^a$ | Evolutionary $^b$ | State |
|----------|--------------|--------------|----|----------------------|--------------|------|----------------|--------|
| 121.30+0.66 | 00 36 47.5   | +63 29 01    | 0.9 | 1.1E+03              | 22           | 1.25 | NRC            |
| 123.07−0.31 | 00 52 25.1   | +56 33 54    | 2.2 | 6.0E+03              | 21.8         | 1.75 | H II           |
| W3(OH)   | 02 27 04.7   | +61 52 26    | 2.4 | 9.5E+04              | 18.1         | 1.50 | UCH II         |
| S231     | 05 39 12.9   | +35 45 54    | 2.0 | 1.3E+04              | 16.4         | 1.50 | NRC            |
| S252A    | 06 08 35.4   | +20 39 03    | 1.5 | 6.4E+03              | 19.2         | 1.75 | H II           |
| RCW142   | 17 50 14.6   | −28 54 32    | 2.0 | 5.7E+04              | 26.1         | 2.25 | UCH II         |
| W28A2    | 18 00 30.4   | −24 03 58    | 2.6 | 2.0E+05              | 22.7         | 2.25 | UCH II         |
| M8E      | 18 04 53.1   | −24 26 42    | 1.8 | 1.7E+04              | 21.9         | 1.75 | UCH II         |
| 8.67−0.36 | 18 06 18.8   | −21 37 38    | 8.5 | 1.3E+05              | 44.2         | 2.00 | UCH II         |
| 10.60−0.40 | 18 10 28.6  | −19 55 49    | 6.5 | 9.2E+05              | 22.4         | 2.50 | CH II          |
| 12.89+0.49 | 18 11 51.1   | −17 31 30    | 3.5 | 3.9E+04              | 23.6         | 2.00 | NRC            |
| W33A     | 18 14 39.2   | −17 52 11    | 4.5 | 1.0E+05              | 20.2         | 1.50 | NRC            |
| 24.49−0.04 | 18 36 05.3  | −07 31 22    | 3.5 | 4.9E+04              | 29.4         | 2.25 | NRC            |
| W43S     | 18 46 03.5   | −02 39 25    | 8.5 | 1.6E+06              | 26.4         | 2.50 | CH II          |
| 31.41+0.31 | 18 47 34.4  | −01 12 46    | 7.9 | 2.3E+05              | 23.1         | 2.25 | UCH II         |
| W44      | 18 53 18.4   | +01 14 57    | 3.7 | 7.5E+05              | 22           | ...  | CH II          |
| 40.50+2.54 | 18 56 10.4  | +07 53 14    | 2.1 | 2.8E+04              | 19.0         | 1.50 |                |
| 35.20−0.74 | 18 58 12.7   | +01 40 36    | 3.3 | 4.5E+04$^d$         | 22           | ...  | H II           |
| 59.78+0.06 | 19 43 11.5  | +23 43 54    | 2.2 | 1.3E+04$^d$         | 22           | ...  | UCH II         |
| ON1      | 20 10 09.1   | +31 31 37    | 6.0 | 1.5E+05              | 20.7         | 1.75 | UCH II         |
| ON2S     | 20 21 41.0   | +37 25 29    | 5.5 | 3.7E+05              | 18.8         | 1.75 | H II           |
| W75N     | 20 38 36.9   | +42 37 38    | 3.0 | 1.4E+05              | 22           | ...  | UCH II         |
| W75OH    | 20 39 00.9   | +42 22 50    | 3.0 | 5.0E+04$^d$         | 22           | ...  | NRC            |
| S140     | 22 19 18.2   | +63 18 48    | 0.9 | 1.9E+04              | 16.9         | 1.25 | NRC            |
| CepA     | 22 56 18.1   | +62 01 47    | 0.7 | 2.2E+04              | 16.9         | 1.50 | UCH II         |
| NGC7538−IRS9 | 23 14 01.6  | +61 27 20    | 2.8 | 3.6E+04              | 18.9         | 1.75 | UCH II         |
| S157     | 23 16 04.3   | +60 01 41    | 2.5 | 2.9E+04              | 15.9         | 0.75 | CH II          |

Notes.

$a$ Power-law index from Mueller et al. (2002).

$b$ NRC: weak or no radio continuum; UCH II: ultra-compact H II; CH II: compact H II; H II: free–free emission from ionized H.

c $T_{iso}$ is taken to be the median value (=22 K) for sources without 350 μm modeling.

d Using $L_{IR}$ in place of $L_{bol}$.

Figure 1. Effective densities as a function of temperature for transitions of CS, HCN, HCO+, N$_2$H+, SO, HNC, and CCH. Collision rates for CCH are assumed to be those of H$_{13}$CN. $n_{eff}$ for SO at $T=10$ K is an upper limit. (A color version of this figure is available in the online journal.)

sample reflects the very heterogeneous populations uncovered in Galactic plane surveys (i.e., BGPS, although this sample is skewed toward more massive/more luminous sources than what is found in BGPS on average) that will be prime candidates for this kind of spectroscopic follow-up.

We describe the observations in Section 2. Integrated intensity maps and correlation between dust emission and molecule emission are explored in Section 3. The physical properties derived from spatial intensity distributions (e.g., size, mass, surface density) are analyzed in Section 4. We summarize our results in Section 5.

### 2. Observations

#### 2.1. Nobeyama Observations

Twenty-seven sources were mapped with BEARS on April 22 through 26, 2004 using the Nobeyama Radio Observatory 45 m telescope. BEARS is a 25 element focal plane array of SIS mixers (Sunada et al. 2000). The beams are spaced at 41.′′1 on the sky with a beamsize of approximately 17.′′8 (see Tatematsu et al. 2004). To map a source at nearly full beam resolution, the array must be dithered in 4×4 grid with a spacing of 20.′′55. Also, to minimize systematic calibration difference between the individual beams, the source was mapped with the array rotated by +90° and −90°. Each source was observed in position-switched mode with OFF positions that were determined to be clean of emission in CS $J=5→4$ (Shirley et al. 2003). The final map covers approximately 3′ × 3′ with 100 spectra.
The BEARS backend is a digital autocorrelator with 1024 channels. The spectral resolution is 37.8 kHz resulting in 0.12 km s\(^{-1}\) resolution at 93.7 GHz. Each spectrum was centered at \(\nu_{LR}\) for the strongest hyperfine component of N\(_2\)H\(^+\) \(J = 1 \rightarrow 0\) \((F\_I = 23 \rightarrow 12\) at 93.1737767 GHz). During data reduction, the spectra were smoothed by four channels to improve the signal-to-noise; the final spectral resolution is 0.49 km s\(^{-1}\).

The spectra were placed onto the \(T_{\ast}^4\) scale by comparing observations of each beam in the array with the Nobeyama S100 receiver. \(T_{\text{mb}}\) is calculated by using a main-beam efficiency of the telescope with S100 of 0.515. The S100 receiver uses the standard chopper-wheel calibration method (Penzias & Burrus 1973). The bright source, G40.50, was observed with each beam in the array on two separate days. The peak of the N\(_2\)H\(^+\) \(J = 1 \rightarrow 0\) emission was used to compute the calibration corrections for each beam. We estimate that the calibration is consistent with \(T_{\ast}^4\) determined by S100 to within 15%. Weather conditions were good to excellent during the observations. The typical wind speed was below 10 m s\(^{-1}\).

The spectra were reduced using the NEWSTAR software package of the Nobeyama Radio Observatory. A linear baseline was removed from each spectrum and the spectra at the same offset position were averaged. Since the actual beam centers on the sky do not form a perfectly regular grid, the spectra were interpolated onto a regular grid using a Gaussian smoothing function. The final spectra were written out as an SPITTS file (Greisen & Harten 1981) and were read into AIPS++ where the individual spectral arrays could be manipulated. Gaussian fits were performed using the profile fitter in AIPS++.

2.2. CSO Observations

Observations of N\(_2\)H\(^+\) \(J = 3 \rightarrow 2\), HCO\(^+\) \(J = 3 \rightarrow 2\), and HCS\(^+\) \(J = 6 \rightarrow 5\) of the 27 sources were performed at the Caltech Submillimeter Observatory 10.4 m telescope on 10 nights between 2002 June and 2004 July. The 230 GHz single-pixel receiver (Kooi et al. 1992, 1998) with a 50 MHz AOS backend was used for all observations. The average spectral resolution was 0.15 km s\(^{-1}\); however, the spectra were smoothed to a velocity resolution of 0.21 km s\(^{-1}\) (N\(_2\)H\(^+\), HCO\(^+\)) and 0.14 km s\(^{-1}\) (HCS\(^+\)). We made oversampled, boustrophedonic (alternating scan direction between left-to-right and right-to-left lines) On-The-Fly maps of N\(_2\)H\(^+\) and HCO\(^+\) on a grid with 10\(^{\prime}\) resolution (Mangum et al. 2007) while HCS\(^+\) observations were single position-switched spectra toward the 350 \(\mu\)m dust continuum peak position.

The standard chopper calibration-wheel method was used to calibrate the spectra on the \(T_{\ast}^4\) scale (Penzias & Burrus 1973). Calibration on planets (Mars, Jupiter, and Venus) was used to convert to the \(T_{\text{mb}}\) scale assuming that the sidelobes do not contribute significantly (Kutner & Ulich 1981). Pointing was performed every hour on planets or bright high-mass star-forming cores (W28A2(1)) when planets were not available. We estimate our pointing is better than 10\(^{\prime}\) when no planets were available and is better than 5\(^{\prime}\) with planets.

2.3. HHT Observations

All 27 sources were mapped using the 1 mm ALMA prototype sideband-separating receiver on the 10 m Heinrich Hertz Telescope (HHT) on Mt. Graham, Arizona over six nights between 2008 October and 2009 June. We observed in dual polarization four IF modes (5.0 GHz IF), with HNC \(J = 3 \rightarrow 2\) (271.9811 GHz) centered in the upper sideband and SO \(N_{\text{J}} = 6_{\text{J}} \rightarrow 5_{\text{J}}\), C\(_2\)H\(_2\) \(N_{\text{J}} = 3_{\text{J}} \rightarrow 2_{\text{J}}\), and CH\(_3\)OH-E \(J_{\text{K}} = 2_{\text{J}} \rightarrow 1_{\text{J}}\) in the lower sideband. We also observed the isotopologues HN\(^{13}\)C \(J = 3 \rightarrow 2\) (261.263 GHz) centered in the upper sideband and \(^{34}\)SO \(N_{\text{J}} = 6_{\text{J}} \rightarrow 5_{\text{J}}\) (256.877 GHz) also centered in the upper sideband between 2009 November and 2010 April. The backend filter banks have 256 channels with 1 MHz width and separation. At these frequencies, we have velocity resolution of ~1 km s\(^{-1}\) and ~30\(^{\prime}\) resolution.

For each source, we made two arcmminute square OTF maps with 10\(^{\prime}\) spacing. All sources were scanned in R.A. and decl. and most were also scanned again in R.A. and decl. with the observational scheme rotated by an arbitrary position angle or offset by 5\(^{\prime}\) from the original map center. Pointing was performed using Jupiter, Saturn, or Mars and the main beam efficiency was calibrated with Jupiter or bootstrapped from the intensity of S140 if Jupiter was not available. Pointing accuracy for the HHT is typically 5\(^{\prime}\) rms.

Spectra were reduced using GILDAS CLASS reduction software. The ALMA prototype receiver is a dual polarization receiver. We scale the horizontal polarization (\(H_{\text{pol}}\)) data to the vertical polarization (\(V_{\text{pol}}\)) data using the main beam efficiency determined for each feed. After a linear baseline was removed, the spectra were convolved onto a 11 \(\times\) 11 10\(^{\prime}\) grid using a Gaussian-tapered Bessel function of the form

\[
\frac{a}{\pi^\frac{1}{4}} \sqrt{\frac{1}{\pi}} \exp \left(-\frac{x^2}{b^2}\right)
\]

where \(a = 1.55(\frac{\nu_{\text{line}}}{\nu_{\text{mb}}})\) and \(b = 2.52(\frac{\nu_{\text{mb}}}{\nu_{\text{pol}}})\) (Mangum et al. 2007). A Bessel function is a more accurate representation of the spatial frequency response of a single dish. Tapering with a Gaussian is a compromise to improve computation time, although the spatial frequencies lost in doing so are not significantly different from those probably already tapered by the response of the receiver feed horns (Mangum et al. 2007).

2.4. Dust Continuum Observations

We also downloaded SCUBA 850 \(\mu\)m archived data (Di Francesco et al. 2008) for the 23 of our sources observed with SCUBA. We display a portion of the image centered on the continuum peak for each available source (see the Appendix).

3. ANALYSIS OF MOLECULAR EMISSION

3.1. Integrated Intensity Maps

We present maps of the integrated intensity of two wavelengths of dust continuum emission and seven molecule transitions. We calculate the integrated intensity from

\[
I(T_{\text{mb}}) = \int T_{\text{mb}} d\nu = \pm \sqrt{\delta v_{\text{line}}} \sqrt{\delta \nu_{\text{chan}}} \sigma T_{\text{mb}},
\]

where \(\delta v_{\text{line}}\) is the full velocity width of the line, the width of the line above \(\sigma T_{\text{mb}}\), and \(\delta \nu_{\text{chan}}\) is the channel width of the spectrometer. Emission maps of integrated intensity, \(I(T_{\text{mb}})\), are shown in the Appendix. All 27 sources are detected in the mapped transitions of HCO\(^+\), N\(_2\)H\(^+\), HNC, CCH, and SO but not every one is detected in the rarer isotopologues (H\(^{13}\)CO\(^+\), H\(^{13}\)NC\(^+\), \(^{34}\)SO) nor the transitions of HCS\(^+\), CH\(_3\)OH, and SO\(_2\). The HCO\(^+\) \(J = 3 \rightarrow 2\) transition typically has the highest integrated intensity with a median intensity, \(I(T_{\text{mb}}) = 61.5\) K km s\(^{-1}\). Figure 2 shows histograms of the integrated intensity for every transition we observed. Observed properties for the mapped transitions are listed in Tables 3–8.

Most transitions trace a single emission peak centered near the dust continuum peak position. The emission maps of N\(_2\)H\(^+\)
are significant exceptions. For 10 of our sources (W3(OH), RCW142, M8E, G8.67, W43S, W44, ON2S, W75(OH), S140, and CepA), we find that N$_2$H$^+$ emission shows significant spatial differentiation that is not reflected in the dust emission or emission from other molecules. These 10 sources not only have multiple intensity peaks in the N$_2$H$^+$ emission, but these peaks are also spatially separated from the peak of the dust and other molecule emission. In general, the dust and other molecules peak near the water maser position while the N$_2$H$^+$ emission in multiple intensity peaks in the N$_2$H$^+$ emission, but these peaks are also spatially separated from the peak of the dust and other molecule emission. These 10 sources not only have differentiation that is not reflected in the dust emission or emission from other molecules. These 10 sources not only have multiple intensity peaks in the N$_2$H$^+$ emission, but these peaks are also spatially separated from the peak of the dust and other molecule emission. In general, the dust and other molecules peak near the water maser position while the N$_2$H$^+$ emission in the differentiated sources peaks several arcseconds ($\sim$15$''$) away from the water maser position — more than can be accounted for with pointing errors alone. Both N$_2$H$^+$ $J = 3 \rightarrow 2$ and $J = 1 \rightarrow 0$ show this differentiation, indicating that this is a spatial variation in the abundance of N$_2$H$^+$ and not a radiative transfer effect.

The spatial variation in the N$_2$H$^+$ emission may be understood from the dependence of N$_2$H$^+$ chemistry on temperature. N$_2$H$^+$ chemistry and HCO$^+$ chemistry are directly related by the gas phase abundance of CO. HCO$^+$ is created by CO while N$_2$H$^+$ is more abundant when CO is frozen out onto dust grains (see Jørgensen et al. 2004). Each source with spatial variations in the N$_2$H$^+$ emission also has centrally peaked HCO$^+$ emission, supporting the conclusion that these sources are chemically differentiated. Warmer temperatures ($T \gg 20$ K) are required to return CO to the gas phase, suggesting that N$_2$H$^+$ is preferentially differentiated toward warmer, more evolved sources.

To test this hypothesis, we looked at several physical properties that might indicate the presence of at least one evolved protostar embedded in the clump. The isothermal dust temperature, $T_{\text{iso}}$, defined as the temperature required to match the mass of the best-fit dust continuum radiative transfer model (Mueller et al. 2002),

$$T_{\text{iso}} = 41(\text{K}) \left[ \ln \left( \frac{M_{\text{iso}}}{5.09 \times 10^{-8} M_{\odot} S_\nu (\text{Jy}) D^2 (\text{pc}) + 1} \right) \right]^{-1} ,$$

should be higher for warmer, more evolved clumps. The spread in $T_{\text{iso}}$ is within a few K of the median $T_{\text{iso}}$ (22 K), reflecting the unresolved temperature gradients found for these sources by Mueller et al. (2002; see Table 2). Warmer clumps will also be more luminous, although more massive clumps will have larger luminosities even at lower temperatures. We make a simple correction for clump mass by taking the ratio of clump luminosity to the dust-determined isothermal mass, $M_{\text{iso}}$.
Table 3  
Observed Source Properties for N$_2$H$^+$ $J = 1 \rightarrow 0$

| Source | $I(T_{	ext{mb}})$ (K km s$^{-1}$) | $T_{	ext{mb}}$ (K) | $\Delta v$ (km s$^{-1}$) | R (pc) |
|--------|----------------------------------|------------------|-----------------|--------|
| S157   | 26.90(2.84)                      | 6.62(0.90)       | 4.4(0.25)       | 0.26(0.04) |
| S252A  | 31.55(3.38)                      | 7.88(0.89)       | 3.8(0.23)       | 0.11(0.020) |
| W3(OH) | 57.50(9.65)                      | 7.30(0.93)       | 5.20(0.25)      | 0.26(0.004) |
| W31(OH)| 24.26(2.05)                      | 7.41(0.54)       | 3.53(0.10)      | 0.14(0.023) |
| W44    | 118.74(14.40)                    | 13.46(1.64)      | 3.20(0.30)      | 0.25(0.044) |
| W28A2  | 119.96(14.42)                    | 14.88(1.82)      | 3.58(0.20)      | 0.18(0.018) |
| M8E    | 36.76(4.46)                      | 7.98(1.00)       | 4.00(0.30)      | 0.17(0.004) |

Note. * M: multiple peaks.

Table 4  
Observed Source Properties for HCO$^+$ $J = 3 \rightarrow 2$

| Source | $I(T_{	ext{mb}})$ (K km s$^{-1}$) | $T_{	ext{mb}}$ (K) | $\Delta v$ (km s$^{-1}$) | R (pc) |
|--------|----------------------------------|------------------|-----------------|--------|
| S157   | 26.90(2.84)                      | 6.62(0.90)       | 4.4(0.25)       | 0.26(0.04) |
| S252A  | 31.55(3.38)                      | 7.88(0.89)       | 3.8(0.23)       | 0.11(0.020) |
| W3(OH) | 57.50(9.65)                      | 7.30(0.93)       | 5.20(0.25)      | 0.26(0.004) |
| W31(OH)| 24.26(2.05)                      | 7.41(0.54)       | 3.53(0.10)      | 0.14(0.023) |
| W44    | 118.74(14.40)                    | 13.46(1.64)      | 3.20(0.30)      | 0.25(0.044) |
| W28A2  | 119.96(14.42)                    | 14.88(1.82)      | 3.58(0.20)      | 0.18(0.018) |
| M8E    | 36.76(4.46)                      | 7.98(1.00)       | 4.00(0.30)      | 0.17(0.004) |

Note. * M: multiple peaks.

Table 5  
Observed Source Properties for N$_2$H$^+$ $J = 3 \rightarrow 2$

| Source | $I(T_{	ext{mb}})$ (K km s$^{-1}$) | $T_{	ext{mb}}$ (K) | $\Delta v$ (km s$^{-1}$) | R (pc) |
|--------|----------------------------------|------------------|-----------------|--------|
| S157   | 26.90(2.84)                      | 6.62(0.90)       | 4.4(0.25)       | 0.26(0.04) |
| S252A  | 31.55(3.38)                      | 7.88(0.89)       | 3.8(0.23)       | 0.11(0.020) |
| W3(OH) | 57.50(9.65)                      | 7.30(0.93)       | 5.20(0.25)      | 0.26(0.004) |
| W31(OH)| 24.26(2.05)                      | 7.41(0.54)       | 3.53(0.10)      | 0.14(0.023) |
| W44    | 118.74(14.40)                    | 13.46(1.64)      | 3.20(0.30)      | 0.25(0.044) |
| W28A2  | 119.96(14.42)                    | 14.88(1.82)      | 3.58(0.20)      | 0.18(0.018) |
| M8E    | 36.76(4.46)                      | 7.98(1.00)       | 4.00(0.30)      | 0.17(0.004) |

Note. * M: multiple peaks.

Table 6  
Observed Source Properties for CCH $N = 3 \rightarrow 2$

| Source | $I(T_{	ext{mb}})$ (K km s$^{-1}$) | $T_{	ext{mb}}$ (K) | $\Delta v$ (km s$^{-1}$) | R (pc) |
|--------|----------------------------------|------------------|-----------------|--------|
| S157   | 26.90(2.84)                      | 6.62(0.90)       | 4.4(0.25)       | 0.26(0.04) |
| S252A  | 31.55(3.38)                      | 7.88(0.89)       | 3.8(0.23)       | 0.11(0.020) |
| W3(OH) | 57.50(9.65)                      | 7.30(0.93)       | 5.20(0.25)      | 0.26(0.004) |
| W31(OH)| 24.26(2.05)                      | 7.41(0.54)       | 3.53(0.10)      | 0.14(0.023) |
| W44    | 118.74(14.40)                    | 13.46(1.64)      | 3.20(0.30)      | 0.25(0.044) |
| W28A2  | 119.96(14.42)                    | 14.88(1.82)      | 3.58(0.20)      | 0.18(0.018) |
| M8E    | 36.76(4.46)                      | 7.98(1.00)       | 4.00(0.30)      | 0.17(0.004) |

Note. * M: multiple peaks.
Table 7

| Source      | I(Tab) (K km s\(^{-1}\)) | T\(_{\text{tab}}\) (K) | Δv (km s\(^{-1}\)) | R (pc) |
|-------------|---------------------------|-------------------------|---------------------|-------|
| 121.30+0.66 | 4.67(0.32)                | 1.10(0.07)              | 2.8(0.17)           | 0.05(0.014) |
| 123.07−6.31 | 6.63(0.38)                | 1.12(0.07)              | 5.3(0.20)           | 0.12(0.018) |
| W3(OH)      | 27.20(1.12)               | 4.03(0.17)              | 5.8(0.06)           | 0.15(0.005) |
| S231        | 6.95(0.49)                | 1.23(0.09)              | 4.6(0.23)           | 0.21(0.024) |
| S252A       | 5.06(0.54)                | 0.99(0.10)              | 2.9(0.19)           | 0.09(0.028) |
| RCW142      | 12.17(1.76)               | 2.72(0.31)              | 5.6(0.79)           | 0.10(0.030) |
| W28A2       | 48.97(2.53)               | 5.29(0.37)              | 8.5(0.47)           | 0.10(0.007) |
| W3E         | 9.94(0.87)                | 2.65(0.22)              | 6.2(0.15)           | 0.07(0.025) |
| W75N        | 36.55(1.11)               | 6.34(0.27)              | 5.1(0.04)           | 0.26(0.006) |
| W75OH       | 35.64(1.18)               | 7.26(0.33)              | 6.0(0.92)           | 0.40(0.018) |
| S140        | 29.29(0.91)               | 8.23(0.33)              | 3.0(0.13)           | 0.12(0.003) |
| CepA        | 31.50(1.00)               | 7.21(0.30)              | 4.7(0.06)           | 0.09(0.000) |
| NGC7538−IRS9| 11.34(0.41)               | 2.71(0.12)              | 3.9(0.11)           | 0.34(0.027) |
| S157        | 6.15(0.48)                | 1.12(0.07)              | 3.5(0.57)           | 0.19(0.059) |

Note. *E: extended emission, no well-defined source boundary.

Table 8

| Source      | I(Tab) (K km s\(^{-1}\)) | T\(_{\text{tab}}\) (K) | Δv (km s\(^{-1}\)) | R (pc) |
|-------------|---------------------------|-------------------------|---------------------|-------|
| 121.30+0.66 | 4.67(0.32)                | 1.10(0.07)              | 2.8(0.17)           | 0.05(0.014) |
| 123.07−6.31 | 6.63(0.38)                | 1.12(0.07)              | 5.3(0.20)           | 0.12(0.018) |
| W3(OH)      | 27.20(1.12)               | 4.03(0.17)              | 5.8(0.06)           | 0.15(0.005) |
| S231        | 6.95(0.49)                | 1.23(0.09)              | 4.6(0.23)           | 0.21(0.024) |
| S252A       | 5.06(0.54)                | 0.99(0.10)              | 2.9(0.19)           | 0.09(0.028) |
| RCW142      | 12.17(1.76)               | 2.72(0.31)              | 5.6(0.79)           | 0.10(0.030) |
| W28A2       | 48.97(2.53)               | 5.29(0.37)              | 8.5(0.47)           | 0.10(0.007) |
| W3E         | 9.94(0.87)                | 2.65(0.22)              | 6.2(0.15)           | 0.07(0.025) |
| W75N        | 36.55(1.11)               | 6.34(0.27)              | 5.1(0.04)           | 0.26(0.006) |
| W75OH       | 35.64(1.18)               | 7.26(0.33)              | 6.0(0.92)           | 0.40(0.018) |
| S140        | 29.29(0.91)               | 8.23(0.33)              | 3.0(0.13)           | 0.12(0.003) |
| CepA        | 31.50(1.00)               | 7.21(0.30)              | 4.7(0.06)           | 0.09(0.000) |
| NGC7538−IRS9| 11.34(0.41)               | 2.71(0.12)              | 3.9(0.11)           | 0.34(0.027) |
| S157        | 6.15(0.48)                | 1.12(0.07)              | 3.5(0.57)           | 0.19(0.059) |

Note. *U: unresolved.
high-mass clumps may generate additional pressure on an individual collapsing core causing the collapse to proceed faster than free fall, thus producing high densities before the molecules responsible for N$_2$H$^+$ destruction (e.g., CO) are frozen out (Lintott et al. 2005). Lintott et al. (2005) suggest that preventing freeze out facilitates N$_2$H$^+$ destruction near massive YSOs and enhances emission from other species that freeze out in more quiescent low mass cores (Myers & Benson 1983; Evans et al. 2001).

Most of our sources do not show large-scale spatial differentiation in any transitions other than the two N$_2$H$^+$ transitions, with the notable exception of ON2S. The 850 $\mu$m continuum emission for ON2S shows two peaks in the larger, more southern source. The northeast peak is brighter than the southwest peak in both 350 $\mu$m and 850 $\mu$m emission indicating that the northwest source is warmer. HCO$^+$, CS, and HNC emission appears to trace the more northern of the two peaks in the 850 $\mu$m continuum, while the position of the N$_2$H$^+$, SO, and CCH emission is more consistent with the position of the more southern peak (see Figure 12 in the Appendix). A colder southern source with strong N$_2$H$^+$ emission is consistent with the predictions for N$_2$H$^+$ chemical differentiation, however the chemical differentiation in SO and CCH is not easily explained. Higher spatial resolution interferometric observations are required to resolve the individual northeastern and southern sources. In a future work, we will use the Submillimeter Array to study the temperature and the chemistry of these sources individually in order to better understand what may be causing the unusual SO spatial differentiation in this source (M. Reiter et al. 2011, in preparation).

Two of the mapped molecules in this survey, CCH and SO, are molecules with unpaired electrons and therefore sensitive to the Zeeman effect and potential tracers of the magnetic fields (e.g., Bel & Leroy 1989, 1998). Zeeman splitting is too small to be observed with our spectral resolution for the transitions considered. The similarity of the CCH and SO emission to the dust continuum emission in this survey indicates that observations of these Zeeman tracers at higher resolution should typically be centered within 15$''$ of the dust continuum position. Ultimately, observing magnetic fields in these clumps via Zeeman splitting in lower frequency transitions of these molecules will best be done with the sensitivity and high spatial resolution available with ALMA (e.g., 3 mm band for CCH and 9 mm band for SO).

3.1.1. Molecular Tracers with Non-detections

HCS$^+$, CH$_3$OH, and SO$_2$, as well as the isotopologues, are detected in some, but not all, of the sources in our sample. In this section we examine the detection statistics and chemistry of HCS$^+$, CH$_3$OH, and SO$_2$.

In the lower sideband of the HNC setup, we have serendipitously detected the CH$_3$OH-E $J_K = 2_1 \rightarrow 1_0$ transition toward 17 clumps (63%). The typical gas phase abundance of CH$_3$OH observed toward high-mass cores cannot be explained by purely gas phase reactions (Leurini et al. 2007). CH$_3$OH is formed in the solid state via ultraviolet irradiation of ices and is released into the gas phase in regions of the core envelope where temperatures are above the gas sublimation temperature of H$_2$O ($T > 90$ K; Charnley et al. 2004). As a result, CH$_3$OH is an excellent tracer of the hot core phase, and the warm emitting regions that are very likely unresolved within our 30$''$ beam. Sources with CH$_3$OH-E $J_K = 2_1 \rightarrow 1_0$ detections have a slightly higher median $L_{bol}$ ($5.7 \pm 4.3 \times 10^4 \ L_\odot$) than the median $L_{bol}$ for the remaining sources ($3.6 \pm 3.0 \times 10^4 \ L_\odot$), suggesting that a hot core is embedded in the clump. With higher spatial resolution, we might expect a better correlation between the intensity of CH$_3$OH and $L_{bol}$. Eleven of the 17 sources with a CH$_3$OH detection are associated with H II emission, although the fraction of the detections with an associated H II region is comparable to the fraction of the total sample with associated H II emission (roughly two-thirds in both cases). This is not surprising given that a protostar with spectral type $\geq$B0 is required to produce enough ionizing radiation to create an observable H II region while a wider range of protostar masses will pass through the hot core phase.

We detect HCS$^+$ $J = 6 \rightarrow 5$ (in the lower sideband for the H$^13$CN $J = 3 \rightarrow 2$ setup at the CSO; Wu & Evans 2003) toward 20 of our 27 sources (74%). The abundance of HCS$^+$ relative to CO is found to be high in cold, dense dark clouds, suggesting that the dissociative recombination of HCS$^+$ to form CS is a slow process, leading to high ratios even in hot ($T \geq 300$ K) environments (Irvine et al. 1983). The seven sources with non-detections also have smaller CS-integrated intensities than those with a detection (median $I(T_{mb}) = 20.2 \pm 2.0$ K km s$^{-1}$ for the non-detections compared to a median $I(T_{mb}) = 44.3 \pm 15.2$ K km s$^{-1}$ for sources with an HCS$^+$ detection). The two other sulfur-bearing molecules we observe also have higher intensities for the sources with an HCS$^+$ detection (SO median $I(T_{mb}) = 9.8 \pm 4.7$ K km s$^{-1}$ for HCS$^+$ detections, compared with $I(T_{mb}) = 6.6 \pm 0.7$ K km s$^{-1}$ for HCS$^+$ non-detections; SO$_2$ median $I(T_{mb}) = 1.4 \pm 0.3$ K km s$^{-1}$ for HCS$^+$ detections, compared to $I(T_{mb}) = 1.1$ K km s$^{-1}$ for CepA, the only SO$_2$ detection without an HCS$^+$ detection). The median $L_{bol}$ for sources with an HCS$^+$ detection is $(9.5 \pm 7.6) \times 10^4$ while it is $(2.2 \pm 0.94) \times 10^4$ for the remaining sources. Sources where we do not detect HCS$^+$ are associated with weak or no radio continuum (two of seven non-detections) or UCH II emission (four of seven non-detections; the remaining source is associated with CH II emission). HCS$^+$ non-detections are preferentially observed toward the earlier evolutionary states.

Our spectral configuration for observing $^{34}$SO included SO$_2$ in the bandpass. We detect SO$_2$ toward 10 of our sources (37%). SO$_2$ is expected in more chemically evolved clumps, as it is created from multiple gas phase reactions of OH with sulfur-bearing molecules (e.g., CS) that are released from grains as protostars heat the envelope (see Section 3.2 for a discussion of SO and SO$_2$ formation). Sources with a SO$_2$ detection tend to have higher $L_{bol}$—half of the sources with a detection have an $L_{bol}$ that is nearly a factor of three larger than the median $L_{bol}$ for sources with no SO$_2$ detection (median $L_{bol}$ for SO$_2$ detections is $(1.0 \pm 0.78) \times 10^5 \ L_\odot$ compared to $(0.36 \pm 0.23) \times 10^5 \ L_\odot$ for the remaining sources). Six of the 10 sources with an SO$_2$ detection are associated with H II emission, again a fraction that is consistent with that observed in the entire sample.

High resolution observations of high-mass star-forming regions reveal many cores in different evolutionary states within a single clump (e.g., Hunter et al. 2008; Brogan et al. 2009), so it may be that a few young cores are responsible for all hot core molecule emission. Many of the sources with a CH$_3$OH or an SO$_2$ detection are also among the nearer sources and more distant sources with a detection are among the more luminous in our sample. This suggests that these transitions are excited in all of the clumps in our sample, but in many cases beam dilution renders the transition unobservable in at 30$''$ beam.
3.2. Integrated Intensity Correlations

We quantify how well the molecules trace the dust (and each other) by calculating the correlation between peak molecular emission and the dust emission at both 350 μm and 850 μm (correlations with 350 μm dust emission are plotted in Figure 3). We use the Bayesian IDL routine LINMIX_ERR (Kelly 2007) to perform a linear regression to find the slope of the best fit line. The relationship between dust emission and molecule emission may not be strictly linear, however, given the scatter in the correlation plots (see Figure 3), we are not justified in fitting a more complex function. We also calculate the Spearman rank coefficient for every combination of dust and molecule intensity. The Spearman rank analysis determines the degree of correlation between two distributions nonparametrically, and is therefore less sensitive to outliers in the data. Larger values of the correlation coefficient ($r_{corr} \geq 0.7$) indicate a more robust correlation.

Correlations of each molecule with the 350 μm and the 850 μm dust emission are in excellent agreement (see Columns 1 and 2, respectively, of Table 9), which is encouraging given that the dust observations have different resolutions (15′′ versus 10′′), and were taken with different telescopes using different observing techniques (jiggle mapping versus scanning). High $n_{eff}$ tracers CCH, HNC, CS, and SO have integrated intensities that are generally well correlated ($r_{corr} \geq 0.7$) with each other. CCH and HNC ($n_{eff} = 1.5 \times 10^5$ cm$^{-3}$ and $n_{eff} = 1.7 \times 10^5$ cm$^{-3}$, respectively) are strongly correlated with every transition. The two densest gas tracers, CS and SO, are poorly correlated with N$_2$H$^+$, again supporting the hypothesis that N$_2$H$^+$ is destroyed in the warm gas in the center of these clumps (see Section 3.1).

CS and SO have somewhat weaker correlation coefficients overall compared to CCH and HNC even though CS and SO are the highest $n_{eff}$ tracers in this survey. This may be due to sulfur’s complicated chemistry, which, in part, remains poorly understood due to the difficulty observing S$^+$ or S II. Beuther et al. (2009) studied the chemistry of sulfur in molecules in their study of high-mass protostellar chemical evolution. As discussed in that work, CS is released from grains early (CS desorbs from grains at temperatures of ~30 K), then reacts with

![Figure 3](image-url)

Figure 3. Correlation between the beam-averaged 350 μm dust intensity and the integrated intensity of the molecules. The Spearman rank correlation coefficients are listed in Table 9. The blue diamonds represent H II regions, magenta triangles represent compact H II regions, and green squares represent UCH II regions. Sources plotted with black crosses have weak or no radio continuum. In some cases, error bars are smaller than the plotting symbol. (A color version of this figure is available in the online journal.)
OH to form SO and eventually SO2. It is possible that shock heating from strong outflows may also drive up gas temperatures, facilitating this chemical pathway. With higher spatial resolution, such as that used in Beuther et al. (2009), we may be able to use the relative abundances of CS, SO, and SO2 to determine the chemical evolutionary state of the individual cores in our sources. Without resolving individual cores, we can say that this kind of chemical evolution likely contributes to the weaker correlations with S-bearing molecules. CS- and SO-integrated intensities are only modestly correlated ($r_{corr} = 0.67$) as we would expect for unresolved chemical differentiation due to the destruction of CS in the production of SO.

Table 9

| Transition | 350 μm | 850 μm | N2H+ (1 → 0) | HCO+ (3 → 2) | N2H+ (3 → 2) | CCH (3 → 2) | HNC (3 → 2) | CS (5 → 4) | SO (67 → 56) |
|------------|--------|--------|---------------|---------------|---------------|------------|------------|------------|-------------|
| N2H+ (1 → 0) | 0.55 | 0.56 | 0.46 | 0.83 | 0.54 | 0.66 | 0.44 | 0.34 |
| HCO+ (3 → 2) | 0.62 | 0.62 | 0.46 | 0.68 | 0.82 | 0.90 | 0.61 | 0.80 |
| N2H+ (3 → 2) | 0.48 | 0.55 | 0.83 | 0.68 | 0.68 | 0.72 | 0.43 | 0.52 |
| CCH (3 → 2) | 0.81 | 0.74 | 0.58 | 0.82 | 0.73 | 0.86 | 0.72 | 0.79 |
| HNC (3 → 2) | 0.78 | 0.73 | 0.68 | 0.90 | 0.76 | 0.86 | 0.68 | 0.76 |
| CS (5 → 4) | 0.75 | 0.70 | 0.44 | 0.61 | 0.43 | 0.72 | 0.68 | 0.67 |
| SO (6 → 5) | 0.68 | 0.73 | 0.34 | 0.80 | 0.54 | 0.79 | 0.76 | 0.67 |

| Number above 0.7 | 7 | 8 | 1 | 4 | 3 | 7 | 6 | 4 | 4 |

Notes. Bold values are greater than the 0.7 threshold to be considered strongly correlated. Imitc values are less than 0.5, the threshold to be considered poorly correlated.

3.3. Column Density Correlations

To compare the behavior of the molecules, we calculate the column density. Molecules with similar $n_{eff}$ will be excited in the same gas, so we expect that as one transition gets brighter, so should other transitions with similar $n_{eff}$, in the absence of strong abundance gradients. Poor correlations may indicate chemical differentiation or the development of substructure (cores within the clump), even if the scales on which these things are occurring are too small to be resolved with the single dish beam.

We calculate the column density from

$$N_{\text{mol}} = \frac{4\pi Q(T_{\text{ex}})e^{E_{m}/kT_{\text{ex}}}}{hc g_{v} A_{\text{mol}}[J_{v}(T_{\text{ex}}) - J_{v}(T_{\text{mb}})]} I(T_{\text{mb}}) \frac{\tau}{1 - e^{-\tau}} \text{ cm}^{-2},$$

where $g_{v}$ is the statistical weight for an energy level $E_{m}$, $A_{\text{mol}}$ is the transition Einstein A, $B_{v}$ is the Planck function, $J_{v}(T) = (h\nu/k)[\exp(h\nu/kT) - 1]$ is the Planck function in temperature units (K), $Q$ is the partition function for the molecule in LTE with excitation temperature $T_{\text{ex}}$, and $\tau/(1 - e^{-\tau})$ is the correction for non-zero optical depth. We approximate the average excitation temperature along the line of sight with the isothermal dust temperature $T_{\text{iso}}$ (see Section 3.1). $T_{\text{iso}}$ averages over significant temperature gradients in the clump, and may not represent the temperature of the gas where a specific transition is excited. However, for the range of $T_{\text{iso}}$ for these clumps, the column density is not very sensitive to the differences in temperature (see Figure 4), changing by at most a factor of two.

We attempt to correct for optical depth, $\tau$, for those molecules for which we have observed a rarer isotopologue. With a value for $\tau$, we can correct the column density for optical depth by multiplying the optically thin column density with a correction factor, $\tau/(1 - e^{-\tau})$ (Goldsmith & Langer 1999). If we assume that the isotopologue ($H^{13}CO^{+}$ for example) is optically thin, then we can use the ratio of the peak temperature of the rarer isotopologue ($H^{13}CO^{+}$) to the peak temperature of the main isotopologue (HCO$^{+}$) to solve for the optical depth. To do so, we assume that the molecular abundance of $H^{13}CO^{+}$ relative to HCO$^{+}$ is equal to the ISM abundance of atomic $^{13}$C relative to $^{12}$C ($^{13}$C/$^{12}$C $= \frac{1}{70}$; Wilson & Rood 1994). With this assumption and ignoring background radiation, we can calculate the optical depth $\tau$ as

$$\frac{[T_{\text{H}^{13}CO^{+}}]}{[T_{\text{HCO}^{+}}]} = \frac{1 - e^{-\tau_{H^{13}CO^{+}}}}{1 - e^{-\tau_{HCO^{+}}}} = \frac{1 - e^{-X\tau}}{1 - e^{-\tau}},$$

where $X$ is the ratio of $^{13}$C/$^{12}$C, $\tau$ is the optical depth of the HCO$^{+}$ line, and $H^{13}CO^{+}$ is assumed to be thin. For HNC, where both isotopologues show line widths that are approximately Gaussian, we find a median $\tau = 3.26 \pm 0.72$. We know that HCO$^{+}$ $J = 3 \to 2$ must be optically thick in at least some of our sources because we see self-absorption in the line profiles. Measuring the temperature of a self-absorbed line is difficult and leads to considerable uncertainty in the temperature of the main isotopologue. ISM abundance ratios for $^{13}$C/$^{12}$C and $^{34}$S/$^{32}$S are also uncertain to at least 10%, further increasing our $\tau$ uncertainty. Not surprisingly, our highly uncertain optical depth correction dramatically increases the scatter in our column densities, and thus decreases the correlation in many cases. We have calculated correlation coefficients for both the optically thin and the optically thick column density and report both in Tables 10 and 11, respectively.

We calculate total H$_2$ column density from the 350 μm dust emission or the 850 μm dust emission for sources without 350 μm data from

$$N_{\text{H}_2} = \frac{S_{\text{obs}}^{\text{H}_2}}{B_{v}(T_{d})\kappa_{v}\mu m_{\text{H}}\Omega_{\text{beam}}},$$

where $S_{\text{obs}}^{\text{H}_2}$ is the observed flux in a 30" aperture and $\kappa_{v} = 0.10 \text{ cm}^2 \text{ g}^{-1}$ is the opacity of gas and dust at 350 μm, and $\kappa_{v} = 0.0197 \text{ cm}^2 \text{ g}^{-1}$ at 850 μm (Ossenkopf & Henning 1994).
Figure 4. Integrated intensity correlations. (a) Chemically opposite species HCO\(^+\) and N\(_2\)H\(^+\) are well correlated (\(r_{\text{corr}} = 0.68\)) inspite of observed spatial differentiation in some sources. (b) Molecules found to trace the dust for low-mass clumps (N\(_2\)H\(^+\)) and high-mass clumps (CS). (c) Nitrogen-bearing molecules that Zinchenko et al. (2009) find to be differentiated. We observe N\(_2\)H\(^+\) differentiation, but our HNC observations do not trace the same peaks as N\(_2\)H\(^+\). (d) The two sulfur-bearing molecules mapped in this study (\(r_{\text{corr}} = 0.67\)). In some cases, the error bars are smaller than the plotting symbol.

Table 10

| Molecule | \(H_2\) | HCO\(^+\) | N\(_2\)H\(^+\) | CCH | HNC | CS | SO |
|----------|--------|----------|----------------|-----|-----|----|----|
| HCO\(^+\) | 0.76   | 0.69     | 0.84           | 0.91| 0.65| 0.77|
| N\(_2\)H\(^+\) | 0.41   | 0.69     | 0.76           | 0.73| 0.45| 0.53|
| CCH     | 0.71   | 0.84     | 0.76           | 0.88| 0.75| 0.70|
| HNC    | 0.71   | 0.91     | 0.73           | 0.88| 0.67| 0.68|
| CS     | 0.62   | 0.65     | 0.45           | 0.75| 0.67| 0.61|
| SO     | 0.78   | 0.77     | 0.53           | 0.70| 0.68| 0.61|

Number above 0.7 4 3 3 5 3 1 2

Notes. Bold values are greater than the 0.7 threshold to be considered strongly correlated. Italic values are less than 0.5, the threshold to be considered poorly correlated.

Table 11

| Molecule | \(H_2\) | HCO\(^+\) | N\(_2\)H\(^+\) | CCH | HNC | CS | SO |
|----------|--------|----------|----------------|-----|-----|----|----|
| HCO\(^+\) | 0.66   | 0.67     | 0.77           | 0.54| 0.35| 0.48|
| N\(_2\)H\(^+\) | 0.41   | 0.67     | 0.76           | 0.39| 0.20| 0.29|
| CCH     | 0.71   | 0.77     | 0.76           | 0.58| 0.54| 0.60|
| HNC    | 0.54   | 0.54     | 0.39           | 0.58| 0.54| 0.51|
| CS     | 0.56   | 0.35     | 0.20           | 0.54| 0.54| 0.50|
| SO     | 0.73   | 0.48     | 0.29           | 0.60| 0.51| 0.50|

Number above 0.7 2 1 1 2 0 0 0

Notes. Bold values are greater than the 0.7 threshold to be considered strongly correlated. Italic values are less than 0.5, the threshold to be considered poorly correlated.

We use \(T_{\text{iso}}\) as the dust temperature and assume a mean molecular weight of \(\mu = 2.34\) as is appropriate for dense molecular gas. We plot the correlation between the molecular column densities and \(H_2\) column density in Figure 5.

Figure 5. Temperature sensitivity of the optically thin column density, assuming local thermodynamic equilibrium.

(A color version of this figure is available in the online journal.)

Correlations between the \(H_2\) column density and the optically thick molecule column densities follow the same general trends as correlations of the molecule-integrated intensities with the dust emission, although this is expected since \(N_{\text{dust}} \sim I(T_{\text{mb}})\) in the optically thin limit. Indeed, the most pronounced differences are for correlations of optical-depth corrected column densities with the \(H_2\) column density. These are made notably worse by the optical depth correction, dropping from strongly correlated (\(r_{\text{corr}} = 0.78\) for the integrated intensity correlation) to relatively poorly correlated (\(r_{\text{corr}} = 0.54\) for the column densities) in the case of HNC.
4. DISCUSSION

With a wealth of galactic plane surveys targeting hitherto unexplored regions of the electromagnetic spectrum coming out in the next few years (e.g., BGPS, ATLASGAL, HIGAL, etc.), it is important to understand how best to follow up the newly discovered sources in these surveys. Molecular line surveys of sources discovered in millimeter and submillimeter dust emission allow us to study the properties of the dense gas in these sources. As discussed in Section 3.1, dense gas tracers such as $N_2H^+$ that are well correlated with the dust for low-mass clumps are not necessarily good tracers of the dust for high-mass sources. Proper assessment of the physical properties of those sources requires an understanding of how the choice of molecular tracer affects the derived source properties. With seven mapped transitions probing more than three orders of magnitude in $n_{\text{eff}}$, we may begin to study how $n_{\text{eff}}$ of the tracer affects the derived physical properties of the clump.

4.1. Size, Mass, and Densities

Our suite of molecular transitions represents a range of excitation conditions, probed by $n_{\text{eff}}$, allowing us to examine how the derived clump properties change as a function of the gas that a specific transition is tracing. In this section, we examine how size $R$, linewidth $\Delta v$, virial mass $M_{\text{vir}}$, mass surface density $\Sigma$, and average density $\langle n \rangle$ vary with $n_{\text{eff}}$ (see Figure 8 and Table 12).

Our sample of high-mass star-forming regions represents a range of masses, chemistries, and evolutionary states. These source-to-source variations are evident in the scatter in the relationships between molecule emission and dust emission (see Sections 3.2 and 3.3 and Figures 3–7). These variations also lead to a broad range of values for the physical properties derived from this sample (though they tend to be positively skewed). We therefore take the median value of each physical property and examine its behavior as a function of $n_{\text{eff}}$ and characterize the spread of values around that median with the median absolute deviation. Because we analyze trends observed in a population of 27 sources, we emphasize that these trends are primarily qualitative as the uncertainties on the median values are large. However, we note that trends in the median value of a given physical property that persist despite the substantial scatter in the values for individual sources are remarkable given the diversity of our sources and relatively small sample size.

4.1.1. Size

Clump size, as determined from a molecular intensity map, depends on the integrated intensity threshold chosen. In this paper, we use the half-peak intensity contour though other definitions that depend on a fixed intensity level are possible. Comparisons between the different definitions generally give the same qualitative conclusions (see Shirley et al. 2003). We find the deconvolved angular radius of a circle with the same...
Figure 7. Column density correlations for the same molecules examined in Figure 4. Note the worsening of the correlation especially for molecules where an optical depth correction has been applied. (a) Chemically opposite species HCO\(^{+}\) and N\(_{2}\)H\(^{+}\). (b) Good dust tracers for low-mass (N\(_{2}\)H\(^{+}\)) and high-mass clumps (CS). (c) Nitrogen-bearing molecules N\(_{2}\)H\(^{+}\) and HNC. (d) The two sulfur-bearing molecules mapped in this study. The modest correlation of the column densities \((r_{corr} = 0.61)\) supports the Beuther et al. (2009) suggestion that SO is formed from CS.

The observed clump size also depends on the underlying density structure of the clump. The large-scale structure of the clumps can be described, on average, with a power-law density distribution \(n \propto r^{-p}\); Mueller et al. 2002; see also Shirley et al. 2003). The convolution of a Gaussian telescope beam pattern with a power-law intensity distribution results in an FWHM core size slightly larger than the beam size with the observed FWHM larger for flatter power-law indices, \(p\) (see Shirley et al. 2003). To account for this, we convolved the N\(_{2}\)H\(^{+}\) \(J = 1 \rightarrow 0\) data to 30\('\) resolution (originally 18\('\)') in order to study all the mapped transitions at the same resolution. With these caveats in mind, we analyze the distribution of sizes for each core in the mapped transitions.

An anti-correlation is observed between the median clump size and the effective density of the tracer (Figure 8). This is the expected trend for a power-law density distribution; the higher density tracers are preferentially excited in a smaller emitting region within the clumps. A similar trend was observed in the HCN and CS multi-transition study of Wu et al. (2010); however, their study contained observations with very different spatial resolutions, from 60\('\) to 30\('\), while the observations in this paper are all similar (30\('\)).

We also see a trend when we look at clump sizes as a function of \(n_{eff}\) for each evolutionary state. If we separate sources into two categories, one for less evolved sources (weak/no radio continuum and UCH\,II) and one for more evolved sources (CH\,II and H\,II), we find that more evolved sources are \(\sim 0.1\) pc larger than less evolved sources, although the difference in size as a function of evolutionary state becomes less pronounced for higher \(n_{eff}\) where source sizes get closer to the beam size. We note that this is a qualitative trend as this analysis requires separating 27 sources into four categories. With so few sources entering into the medians, the median absolute deviation is dominated by small number statistics.

### 4.1.2. Masses

Assuming these clumps are self-gravitating, we calculate the virial mass from

\[
M_{vir} = \frac{5R\Delta v^{2}}{8a_{1}a_{2}G \ln(2)} \approx 209 M_{\odot} \left(\frac{R/1\,\text{pc}}{\Delta v/1\,\text{km}\,\text{s}^{-1}}\right)^{2} a_{1}a_{2}
\]

\[
a_{1} = \frac{1 - p/3}{1 - 2p/5}, \quad p < 2.5,
\]

where \(a_{1}\) is the correction for power-law density distribution and \(a_{2}\) is the correction for non-spherical shape (Bertoldi & McKee 1998).
For aspect ratios less than 2, $a_2 \approx 1$, so we do not include its effects here. We use power-law indices $p$ found in Mueller et al. (2002) to calculate $a_1$. We assume a value of $a_1 = 1.39$ (corresponding to the median power-law index, $p = 1.75$) for sources without 350 μm modeling.

In the $M_{\text{vir}}$ calculation, the linewidth, $\Delta v^2$, reflects the thermal energy and turbulent pressure supporting the cloud against gravity. Linewidths may be broadened by optical depth and systematic collapse motions (for example), leading to an artificially large estimate of the mass. Observing a rarer isotopologue of the molecule mitigates the effects of optical depth in some cases, as the optically thin linewidth from the rarer isotopologue may be substituted for the optically thick linewidth of the main isotopologue. Other systematic motions can still broaden the linewidth and, in many cases, we cannot correct for those effects. Thus, any mass estimate dependent on the linewidth is only an upper bound.

We have observed isotopologues of four species in this survey, H13CO+, HN13C, C34S, and 34SO, and have derived reliable linewidths for the first three. Comparing the median $M_{\text{vir}}$ calculated with the main and the rarer isotopologue linewidths, we find that the broader linewidths of the main isotopologues appear to be ∼0.1 pc larger on average.

![Figure 8](image1)

**Figure 8.** As a function of $n_{\text{eff}}$ (at 20 K), we show (a) median clump size, (b) median $\Sigma$, (c) median $\langle n \rangle$, (d) median $M_{\text{vir}}$, (e) median $\Delta v$, and (f) median $L'$ for each tracer. Median clump size and $L'$ appear to be anti-correlated with $n_{\text{eff}}$ while $\Sigma$ and $\langle n \rangle$ show a positive correlation. $M_{\text{vir}}$ and $\Delta v$ do not appear to be correlated with $n_{\text{eff}}$. N2H$^+$ $J = 1 \rightarrow 0$ observations have been convolved to 30'' resolution to match the resolution of the other transitions.

![Figure 9](image2)

**Figure 9.** Median clump size for each evolutionary state as a function of tracer $n_{\text{eff}}$ (at 20 K). Points for each evolutionary state are plotted near the tracer $n_{\text{eff}}$ and gray boxes indicate points belonging to a single tracer. Younger evolutionary states (weak or no radio continuum—"no RC," UCHII) are denoted with black crosses and green squares, respectively. More evolved stages (CHII and HII) are denoted with magenta triangles and blue diamonds, respectively. More evolved clumps appear to be ∼0.1 pc larger on average.
lead to a systematic overestimation of the mass by a factor of up to 1.83 (see Figure 9).

$\Delta v$ is also not correlated with $n_{\text{eff}}$ (see Figure 8). For every source with a reliable fit for the linewidth, the optically thin isotopologue has a smaller linewidth than the main isotopologue, although this difference is most pronounced for HCO$^+$, which we expect to be optically thick (see Figure 10). No correlation emerges when we compare the optically thin isotopologues with $n_{\text{eff}}$, we find only that the scatter in the lack of correlation is somewhat decreased.

### 4.1.3. Mass Surface Density and Volume Density

McKee & Tan (2003) first noted that almost all regions forming massive stars have mass surface densities $\Sigma \sim 1$ g cm$^{-2}$ (equivalent to a column density of $>2 \times 10^{21}$ cm$^{-2}$). Recent numerical simulations by Krumholz et al. (2007a) and Krumholz & McKee (2008) have shown that $\Sigma \geqslant 1$ g cm$^{-2}$ is a minimum threshold to prevent fragmentation of a massive clump into several smaller objects. $\Sigma$ of order 1 corresponds to an opacity high enough for gas to be heated substantially (hundreds of K) and to substantial distances ($\geqslant 1000$ AU) from the protostar (Krumholz et al. 2007a). Warmer temperatures are crucial to prevent the clump from fragmenting into smaller objects, as the Jeans Mass decreases for higher densities for a fixed temperature ($M_J \sim T^{3/2}$).

We calculate the mass surface density of our sources using

$$\Sigma = \frac{M_{\text{vir}}}{\pi R^2} \approx 0.665 \frac{(M_{\text{vir}}/1.0 \times 10^4 M_\odot)}{(R/1 \text{ pc})^2} \text{ g cm}^{-2}. \quad (10)$$

The median surface density for each transition is positively correlated with $n_{\text{eff}}$. This is as expected; higher $n_{\text{eff}}$ transitions trace denser regions of the clump, corresponding to larger surface densities. Surface densities are $\geqslant 1$ for $\sim 1/3$ of our sources when calculated from transitions with $n_{\text{eff}} \approx 10^4$ cm$^{-3}$. When measured with a tracers with $n_{\text{eff}}$ above a few $10^6$ cm$^{-3}$, $\sim 2/3$ of the clumps have $\Sigma \geqslant 1.0$ g cm$^{-2}$. Unlike median size, median surface density shows no apparent trend with evolutionary state.

Krumholz & McKee (2008) point out that even though suppression of fragmentation allows massive stars to form from massive cores, some low-mass stars will form before radiation halts fragmentation. Therefore, massive stars are expected to be found in clusters. Volume-averaged density, $(n)$, for a clump with clustered star formation embedded within will be dominated by the diffuse gas between the high density cores, leading to smaller values of $(n)$ that would be expected for an
individual star-forming core (Dunham et al. 2010). \((n)\) is then a valuable tool for determining whether objects discovered in new and upcoming surveys (e.g., BGPS) are individual star-forming cores or clumps with cores embedded within. For this survey, all of our sources are clumps rather than individual cores, so we focus on characterizing the variation in \((n)\) with \(n_{\text{eff}}\).

Assuming these clumps are spherical, we can calculate the volume-averaged density,

\[
\langle n \rangle = \frac{M_{\text{cl} \text{eff}}}{\frac{4}{3} \pi R^3 \mu m_H},
\]

where \(\mu\) is the mean molecular weight of molecular gas (\(\mu = 2.34\)), \(m_H\) is the mass of hydrogen, and \(R\) is the radius of the half-peak intensity contour. We find that \((n)\) increases with \(n_{\text{eff}}\), exactly as we expect for transitions with higher \(n_{\text{eff}}\) tracing denser gas (see Figure 7). As a function of evolutionary state, we find that the \((n)\) values for sources associated with no \(H\ \Pi\) emission or (U)CH \(\Pi\) emission are essentially the same, while \((n)\) for sources associated with \(H\ \Pi\) regions tends to be higher. The difference is small and may simply reflect small number statistics.

### 4.2. Comparison of Molecular Column Density with Physical Variables

We find poor correlations (\(r_{\text{corr}} < 0.7\)) between column density and the physical variables \(L_{\text{bol}}, M_{\text{iso}}, \langle n \rangle, \) and \(\Sigma\). Several factors contribute to these poor correlations. A specific molecular transition will trace out the parts of a clump with conditions sufficient to excite that transition. Clumps likely contain unresolved cores that, if they are forming stars, have strong gradients with temperature and density increasing toward the center of the core. Chemical variability from source to source as well as within our beam increases the scatter in comparisons of the column densities.

Qualitatively, the column density of higher \(n_{\text{eff}}\) tracers tend to be larger for larger \(L_{\text{bol}}\) and \(M_{\text{iso}}\), though the correlation never becomes robust (\(r_{\text{corr}} > 0.7\)). Both volume-averaged density, \((n)\), and mass surface density, \(\Sigma\), are completely uncorrelated with molecular column densities regardless of \(n_{\text{eff}}\). Ratios of column densities also fail to correlate with any of the physical variables. As discussed in the previous section, the various tracers in this study are sensitive to different excitation conditions and we believe that analyzing the median value of the clump properties is the best way to understand how derived clump properties depend on the tracer chosen.

### 4.3. \(L'_{\text{mol}}\) Comparisons

Understanding clustered, massive star formation in our own galaxy is an essential step to understanding star formation in other galaxies. Except for the nearest galaxies, star formation can only be assessed via the unresolved star formation indicators across a galaxy. Beginning with Schmidt (1959), star formation prescriptions in other galaxies have relied on the correlation between the galaxy’s gas content and its star formation rate (SFR). Kennicutt (1998) refined this relationship using \(H\ I\) and CO emission to measure the amount of atomic and molecular gas available for star formation, yielding a power-law index of 1.4–1.5. Molecular gas is essential to star formation, but much of the molecular gas traced by CO (\(n_{\text{crit}} = 2.8 \times 10^5\) cm\(^{-3}\), \(n_{\text{eff}} = 30–50\) cm\(^{-3}\)) is probably diffuse, and therefore may not be directly involved in the star formation process (Gao & Solomon 2004).

More recent work has focused on higher density tracers such as HCN \(J = 1 \rightarrow 0\) (\(n_{\text{crit}} = 1.2 \times 10^6\) cm\(^{-3}\), \(n_{\text{eff}} = 1.4 \times 10^4\) cm\(^{-3}\)) and HCN \(J = 3 \rightarrow 2\) (\(n_{\text{crit}} = 6.8 \times 10^7\) cm\(^{-3}\), \(n_{\text{eff}} = 1.5 \times 10^5\) cm\(^{-3}\)) to study the dense molecular gas that is much more likely to be associated with star formation (Gao & Solomon 2004; Bussmann et al. 2008). The star formation law derived from studying dense molecular gas yields a power-law index much closer to unity. Wu et al. (2010) extended the HCN \(J = 1 \rightarrow 0\) sample to galactic cloud cores and found that the data were still well fit with a power-law index of 1. This led to the interpretation that the linear correlation between dense molecular gas and the SFR may be understood as the cumulative contribution of individual star-forming “units.”

Extragalactic star formation studies have used the molecular line luminosity, \(L'\), to estimate the mass of dense gas available for star formation (see Juneau et al. 2009 for a summary of dense gas tracers). \(L'\) is defined as

\[
L' = 23.504\pi \Omega_{\ast\ast}\beta (1 + z)^{-3} D_L^2 I(T_{\text{mb}}) \quad (12)
\]

where \(D_L\) is the luminosity distance in Mpc, \(z\) is the redshift, and \(\Omega_{\ast\ast}\) is the solid angle of the source convolved with the beam in arcsec\(^2\). As described in Mangum et al. (2008), two lines with the same brightness temperature and spatial extent will have the same \(L'\), such that ratios of \(L'\) indicate the physical conditions in the gas.

Empirically, \(L'\) is found to trace the mass of molecular gas,

\[
M_{\text{dense}} = \alpha L'_{\text{mol}},
\]

where \(\alpha\) is a scaling factor that has traditionally been based on the \(H_2\)-mass-to-line-luminosity found in the Milky Way (Downes et al. 1993). A direct test of how well \(L'\) actually traces the mass of dense gas is not possible without an independent determination of the mass. For our galactic clumps, we use \(M_{\text{iso}}\) (see Section 3.1; Mueller et al. 2002), which is determined from the optically thin 350 \(\mu\)m dust flux in a 120” aperture. \(M_{\text{iso}}\) is independent of the integrated intensity of line and independent of the source size determined from the map of that line, making this method of determining mass completely independent from \(L'\).

Comparing \(L'\) with \(M_{\text{iso}}\), we find that the two are highly correlated (\(r > 0.7\); see Figure 11) for every transition. Using LINMIX_ERR to do a linear regression, we find power-law indices that are slightly less than 1 in every case. There is no trend in power-law index with \(n_{\text{eff}}\). Even \(N_2H^+\), which shows strong chemical differentiation in the integrated intensity maps and the poorest correlations with the dust flux, has an \(L'\) that is well correlated with \(M_{\text{iso}}\) (with somewhat more scatter than the other transitions). All of the dense gas tracers we have observed, despite chemical and kinematical behavior that affects the other derived properties of these clumps, are linearly related to the mass of dense gas available for star formation. Power-law indices slightly smaller than 1 indicate that \(L'\) will tend to underestimate the mass, and we show that using higher \(n_{\text{eff}}\) tracers will not resolve this issue.

Looking at \(L'\) versus \(M_{\text{vir}}\) for multiple transitions of HCN and CS, Wu et al. (2010) also found that \(L'\) is an excellent tracer of the dense gas. As those authors discuss, questions have been raised concerning how reliably HCN, CS, and other high \(n_{\text{eff}}\) molecules trace the mass, given the high densities required for excitation. The power indices near unity we find support
the Wu et al. (2010) conclusion that $L'$ is linearly related to the mass regardless of $n_{\text{eff}}$ of the tracers in these two studies because they trace only the dense gas most relevant to star formation.

5. CONCLUSIONS

We have observed 27 sources in 12 transitions of eight molecules spanning three orders of magnitude in $n_{\text{eff}}$. These transitions are observed with nearly identical resolution ($\sim 30''$), allowing for inter-comparison of the clump properties derived from the mapped transitions. Understanding how derived clump properties depend on $n_{\text{eff}}$ of the tracer is essential for designing and interpreting molecular line follow-up observations of sources discovered in continuum surveys of the Galactic plane.

We confirm the Zinchenko et al. (2009) result that N$_2$H$^+$ does a poor job of tracing the dust emission in our massive star-forming clumps. This is evident in the spatial differentiation of N$_2$H$^+$ emission in both the $J = 1 \rightarrow 0$ and $J = 3 \rightarrow 2$ transitions, and the weak correlations between N$_2$H$^+$ and dust emission. N$_2$H$^+$ is a better tracer of the dust in early stages of evolution, but for our more evolved clumps we find that clump luminosity and associated H II emission support the conclusion that N$_2$H$^+$ is differentiated in these sources because they are warmer and thus CO, which destroys N$_2$H$^+$, is in the gas phase.

Unlike N$_2$H$^+$, high $n_{\text{eff}}$ molecules are well correlated with the dust emission and each other. The correlation between CS and SO column densities is somewhat lower ($r_{\text{corr}} = 0.61$), as we would expect if CS is destroyed in the production of SO, as suggested by Beuther et al. (2009).

Derived clump properties vary with $n_{\text{eff}}$ as we would expect. The highest $n_{\text{eff}}$ tracers are sensitive to dense gas in the center of these clumps, corresponding to a smaller emitting region (size) but a higher surface density. Virial masses should not depend on $n_{\text{eff}}$ of the tracer because the calculation is dominated by $\Delta v^2$, which represents thermal and turbulent support in virialized gas. We note that virial mass is particularly sensitive to optical depth effects in the line, leading to an overestimate of the mass by almost a factor of two if an optically thin isotopologue is not used to determine $\Delta v$.

Comparing median clump properties—size $R$, mass $M_{\text{vir}}$, surface density $\Sigma$, and volume density $\langle n \rangle$—to $n_{\text{eff}}$, we find larger median sizes for lower $n_{\text{eff}}$ transitions. More evolved sources appear to be ~0.1 pc larger than younger sources. Virial masses derived from optically thick lines may not depend directly on $n_{\text{eff}}$ since these linewidths may be significantly broadened by non-virial motions. Such effects could lead to possible over- (or under-) estimates of $M_{\text{vir}}$. Virial masses derived from more optically thin lines trace greater pathlengths through the clumps and may therefore be more reliable. Mass surface density is correlated with $n_{\text{eff}}$, as we expect from virial masses that are not correlated with $n_{\text{eff}}$ and sizes that decrease with increasing $n_{\text{eff}}$. Approximately $\frac{3}{4}$ of our sources have $\Sigma \geq 1 \, \text{g cm}^{-2}$ (as measured by tracers with $n_{\text{eff}} > 10^4$) which Krumholz & McKee (2008) identify as a minimum surface density required to suppress fragmentation. Volume density is correlated with $n_{\text{eff}}$ and may be useful for distinguishing between higher density cores (dominated by the high column density of the forming protostar) and lower density clumps (dominated by the diffuse gas between cores embedded in the clump). Molecule column densities are not well correlated with any of the physical properties we derive for these clumps.

We find a nearly linear relationship between $L'$ and $M_{\text{iso}}$, independent of $n_{\text{eff}}$ of the tracer, confirming that $L'$ does reliably trace the mass of dense molecular gas.

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APPENDIX

APPENDIX MATERIAL

Contour maps are presented for the 27 high-mass star-forming regions in Figure 12. Beamsizes are indicated in the lower left hand corner of each panel. The lowest contour as a percentage of the peak intensity and the corresponding $\sigma$ value are listed at the bottom of each panel. Contours are in steps of percentages of the peak. In cases where a percentage less than that of the lowest contour were used, step size is indicated in percent of the peak after the $\sigma$ value.

Source characteristics are listed in Table 2.
Figure 12. Contours of the dust continuum intensity and molecular integrated intensity observed toward each source in the survey. Positional offsets are indicated in arcseconds from the center position listed in Table 2. Beamsizes are indicated in the lower left hand corner of each panel. The lowest contour as a percentage of the peak intensity and the corresponding $\sigma$ value are listed at the bottom of each panel. Contours are in steps of percentages of the peak. In cases where a percentage less than that of the lowest contour were used, step size is indicated in percent of the peak after the $\sigma$ value.

(A extended version of this figure is available in the online journal.)

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