CLOUDS AND CHEMISTRY IN THE ATMOSPHERE OF EXTRASOLAR PLANET HR8799b

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ABSTRACT

Using the integral field spectrograph OSIRIS, on the Keck II telescope, broad near-infrared H- and K-band spectra of the young exoplanet HR8799b have been obtained. In addition, six new narrowband photometric measurements have been taken across the H and K bands. These data are combined with previously published photometry for an analysis of the planet’s atmospheric properties. Thick photospheric dust cloud opacity is invoked to explain the planet’s red near-IR colors and relatively smooth near-IR spectrum. Strong water absorption is detected, indicating a hydrogen-rich atmosphere. Only weak CH4 absorption is detected at K band, indicating efficient vertical mixing and a disequilibrium CO/CH4 ratio at photospheric depths. The H-band spectrum has a distinct triangular shape consistent with low surface gravity. New giant planet atmosphere models are compared to these data with best-fitting bulk parameters, T_{\text{eff}} = 1100 K ± 100 and log(g) = 3.5 ± 0.5 (for solar composition). Given the observed luminosity (log L_{\text{obs}}/L_{\odot} ≈ −5.1), these values correspond to a radius of 0.75 R_{\text{Jup}} ± 0.12 and a mass of ∼0.72 M_{\text{Jup}} ± 0.05—strikingly inconsistent with interior/evolution models. Enhanced metallicity (up to ∼10× that of the Sun) along with thick clouds and non-equilibrium chemistry are likely required to reproduce the complete ensemble of spectroscopic and photometric data and the low effective temperatures (<1000 K) required by the evolution models.

Key words: brown dwarfs – planetary systems – stars: atmospheres – stars: low-mass

Online-only material: color figures

1. INTRODUCTION

Direct imaging of exoplanets is now becoming a regular occurrence with the recent discoveries of 2M1207b (Chauvin et al. 2005), Fomalhaut b (Kalas et al. 2008), Beta Pic b (Lagrange et al. 2009), 1RXS J1609 b (Lafrenière et al. 2008), and the quadruple-planetary system HR8799b, c, d, and e (Marois et al. 2008, 2010). HR8799 is, at present, a rare find among exoplanets with its four potentially massive planets (∼5–10 M_{\text{Jup}}) orbiting a ∼30 Myr A5 star (Zuckerman et al. 2011) and is the first and, so far, only multi-planet system to be directly imaged. All four planets are accessible to photometric and spectroscopic follow-up observations across many wavelengths and, in this paper, near-IR spectra of HR8799b are folded into a large set of photometric data for a broadband analysis of the planet’s photospheric properties.

Since its discovery, a number of new photometric and astrometric measurements of the HR8799 system have been made. These new data have broadened the photometric wavelength coverage and extended the time coverage of orbital motion (Metchev et al. 2009; Lafrenière et al. 2009; Hinz et al. 2010; Currie et al. 2011). The very precise astrometric data have already inspired a number of papers focused on the orbital dynamics and formation of the system; however, the conclusions about the planetary masses and long-term stability of the system from these works are mixed (Fabrycky & Murray-Clay 2010; Reidemeister et al. 2009; Godziewski & Migaszewski 2009; Dodson-Robinson et al. 2009; Kratter et al. 2010). As this paper focuses on photospheric properties, the dynamical nature of the system will not be discussed here.

The goal of this paper is to characterize the properties of HR8799b, independent of what interior/evolution models predict and independent of what the age or dynamics of the system suggest. Following a description of the new spectra and photometry added by this work (Section 2), the available set of photometric and spectroscopic data are compared to similar data for brown dwarfs (BDs)—the closest analogs to young giant planets (Section 3). The atmospheric cloud and chemical properties of HR8799b are inferred by comparing a new set of substellar atmosphere models to the available set of photometry and spectroscopy (Sections 4 and 5). The results are compared to the predictions of interior/evolution models. Also, given the potential for similarities, a comparison is made between HR8799b and 2M1207b (Section 6). The conclusions from this work are summarized in the last section.

2. OBSERVATIONS AND DATA REDUCTION

2.1. OSIRIS Spectroscopy

Observations of HR8799b were made on 2009 July 22, 23, and 30 (UT) and on 2010 July 11 and 13 (UT) at the W. M. Keck Observatory. Natural guide star adaptive optics and the OH-Suppressing Infra-Red Imaging Spectrograph (OSIRIS) instrument (Larkin et al. 2006) were used on the Keck II telescope. Broadband H and K spectra were obtained over a 0′′.32 × 1′′28 patch of the sky surrounding the planet at 0′′.02 spaxel−1 (see Figure 1), where spaxel refers to one OSIRIS spatial resolution element. OSIRIS is an integral field spectrograph that uses a grid of lenslets in the focal plane to dissect the field of view (FOV). The light from each lenslet is dispersed by a grating, producing an array of spectra that are folded into a large set of photometric data for a broadband analysis of the planet’s photospheric properties.

The OSIRIS DRP was used to produce basic calibrated data (BCD) and includes spatial rectification of the raw data,
wavelength and dispersion solutions, sky subtraction, cosmic ray rejection, and so forth (Krabbe et al. 2004). On all three nights in 2009, OSIRIS was operating above its normal detector temperature (~68 K). On July 22, the temperature was ~79 K, but holding steady. On July 23, the temperature was ~80 K and also stable throughout the night. However, on July 30, the temperature was not stable and increased from 83.56 K at the beginning of the observations to 84.54 K at the end. During this period of unstable temperatures, the detector noise in the dark frames increased exponentially. To avoid taking 900 s sky and dark frames repeatedly throughout the night, “scaled darks” were generated for each 900 s exposure on the science target by scaling actual darks taken at the beginning and end of the half-night. An exponential function was fit to the total counts in each actual dark from all three nights and, using this function, scaled darks were generated for the science-target cubes on July 30.

In 2010, OSIRIS was functioning perfectly, and the observing conditions were excellent.

The final BCD data cubes were, at every spatial position, divided by the spectrum of an A0 standard star and multiplied by a blackbody (T = 9600 K). The spectrum of the A0 star was extracted from its data cube by fitting a two-dimensional Gaussian at every monochromatic image slice, rather than using the fixed-aperture extraction routine in the OSIRIS DRP. BD+14 4774 was discovered to be a binary (0.07 separation) with a likely mid-to-late-type companion; for this star, a two-component Gaussian point-spread function (PSF) was used to extract the A0 fluxes. The final BCD cubes are cubes with wavelength and two spatial dimensions. At both H and K bands, the planet was clearly visible in all median-collapsed cubes (median along wavelength). The nominal resolving power of OSIRIS is R ~ 4000 with 1600 wavelength channels. However, to improve the signal-to-noise ratio (S/N), the BCD cubes were binned to 25 wavelength channels, down to R ~ 60. The wavelength positions of OH skylines and telluric star features were measured and no artificial wavelength shifts were discovered; a potential concern given the temperature variations encountered during the observations.

While the planet is easily detected in all of the BCD cubes, contamination by scattered starlight remains a concern. Figure 1 shows the FOV and orientation for the OSIRIS observations and the extent of speckle contamination in the observed region. Since the spatial scaling of the speckles, radially from the star, is proportional to λ, speckles can intersect the spatial location of the planet in a wavelength-dependent manner. To suppress the speckle pattern, each BCD cube was processed with custom IDL routines that first subtracts a (spatially) low-pass filtered version of the image slice in each wavelength channel and rebins the data cube (by taking the median in the λ dimension) to 25 wavelength channels (R ~ 60). Each monochromatic image slice of this rebinned cube is magnified about the stars position by λm/λ, where λm is the median wavelength in either H or K band. In these shifted data cubes, individual speckles align and are fit with simple polynomial functions of λ. Prior to this step, the planet is masked to avoid biasing the polynomial fit. The final polynomial fits to the speckle fluxes are subtracted from the shifted cube before finally reversing the magnification of the cube, placing the planet back in its original position. Collapsed H-band data cubes, before and after speckle removal, are shown in Figure 2 (a similar improvement is seen in the before and after K-band data cubes). The speckle-suppression scheme described here is similar to that used for the low-mass companion GQ Lupi b (McElwain et al. 2007); however, McElwain et al. did not rescale the cubes for speckle fitting but instead fit the underlying speckle halo in each spectral bin of the data cube since each spatial location is more than critically sampled.

The speckle-suppression steps described above contain two important variables: the size of the masked region (centered on the planet) and the order of the polynomial used to fit the speckle fluxes. To determine the best pair of parameters, 10 fake planets (similar in brightness and FWHM as the real planet but with constant Fν) were inserted into the BCD data cubes, evenly spaced but avoiding the real planet location as well as the edges of the FOV. These cubes were processed using the 12 combinations of four mask radius values (from 0.5 to 2 in 0.5 spaxel steps) and three polynomial orders (1, 2, and 3). The parameters resulting in extracted fake planet spectra with the smallest mean deviation from a flat line and smallest rms(λ) were

4 Several updates of the OSIRIS DRP became available during the course of this work; one of these was directly related to preventing unwanted wavelength shifts (http://irlab.astro.ucla.edu/osiris/).
mask radii equal to 1.5 and 2.0 for $H$ and $K$ bands, respectively and a polynomial order equal to 1 for both bands. The mean residuals of all 10 fake planet spectra are shown in Figure 3, straddling zero flux. The rms($\lambda$) for the fake planet spectra were comparable to, or smaller than, the final uncertainties adopted for the planet spectrum, providing confidence that the errors have not been significantly underestimated. Furthermore, the data collected in 2009 compare very well to those from 2010, indicating that the thermal instabilities and scaled darks had little impact on the final reduced spectra.

The planet fluxes were obtained from the speckle-suppressed data cube by convolving each monochromatic image slice with a two-dimensional, Gaussian PSF with FWHM = 2.2 ($H$-band) and 2.7 ($K$-band) spaxels. The final spectra were flux calibrated using $H$ band (Metchev et al. 2009) and revised $K_s$ magnitudes (discussed below). The final extracted spectra and corresponding error bars (taken as the rms of all the extracted spectra in each bandpass) are shown in Figure 3 and listed in Table 2.

As can be seen from the residuals of the fake planet fluxes (Figure 3), the noise in the spectra is correlated between different wavelengths. This is a natural consequence of the fact that the noise is primarily from the speckle pattern—the typical speckle has a monochromatic size of $\lambda/D$, as does the extraction box for the planetary spectrum. A single speckle’s position scales with wavelength proportional to $\lambda$. At $K$ band, the planet is located approximately 42 $\lambda/D$ from the star. For the speckle to move completely across the extraction box requires a change in wavelength of $\frac{(43/42-1)}{2.1 \mu m} = 0.05 \mu m$. Hence, speckle noise patterns will be correlated across the spectrum over a range of 0.05 $\mu m$ (Sparks & Ford 2002). In principle, this should be formally accounted for in the estimate of $\chi^2$ when fitting models to the data (described below). However, since the spectral features being fit are also of width $\sim 0.05 \mu m$ or larger, the correlations in the noise have been ignored.

The planet spectra, before and after speckle suppression, are also compared in Figure 3. Near 1.7 $\mu m$, the impact of a bright speckle can be seen as a fairly sharp peak spanning two wavelength channels. Speckle noise and photon noise are roughly equal in these data sets at this binning, so speckle suppression improved the S/N by approximately 50%. Most importantly, the speckle noise that was removed was spectrally correlated across multiple channels and non-Gaussian, e.g., the feature at 1.7 $\mu m$ in Figure 3, and would have a much more significant effect on the final interpretation than the white photon noise.

2.2. Photometry

On 2009 August 7 (UT), when OSIRIS was too warm to operate, an angular differential imaging (ADI; Marois et al. 2006) sequence was taken with NIRC2 using four narrowband filters in the $K$-band window (see Table 3). These filters were selected to sample, as uniformly as possible, the full $K$-band region accessible with OSIRIS and thus serve as a consistency check on the shape of the final, fully processed, OSIRIS spectrum. These narrowband data are compared in Figure 3 to equivalent flux points obtained from the OSIRIS data convolved with the filter response profiles. The NIRC2 narrowband colors are consistent with the OSIRIS spectral shape on either side of the peak flux; however, HeIB (2.06 $\mu m$) photometry is only consistent at $\sim 2\sigma$. The overall agreement, between these two non-contemporaneous data sets, processed with different reduction software, provides some confidence that the planet’s spectrum was faithfully extracted from the speckle-contaminated data.

Also in 2009, new CH$_4$ (4% S/L) photometry was obtained with the Near-Infrared Coronagraphic Imager (NICI; Toomey & Ftaclas 2003) on Gemini South. The NICI CH$_4$ short/long filters more optimally probe the strength of CH$_4$ absorption than the NIRC2 CH$_4$ short/long filters. The former are narrower, have less overlap, and have central wavelengths that correspond well to the min and max fluxes seen in $H$-band spectra of mid to late T dwarfs. A comparison between these new photometric data and equivalent band-integrated flux points from the OSIRIS spectrum is shown in Figure 3. Here again the CH$_4$ short/long slope is in excellent agreement with the OSIRIS spectrum.

Since the discovery of the HR8799 planetary system more broadband observations have been published and improvements have been made in the Locally Optimized Combination of Images (Lafrenière et al. 2007) and ADI algorithms (at least those used by this group). For the analysis presented here, the $H$-band photometry of Metchev et al. (2009) has been adopted. Also, $K_s$ data taken in 2010 have been analyzed and a new $K_s$ magnitude ($M_{K_s} = 14.15 \pm 0.1$) has been obtained. See Table 3 for a full list of the NIRC2 and NICI photometry used in this study.

2.3. Comparison to Kn3 Spectrum

On 2009 June 21 (UT), Bowler et al. (2010) observed HR8799b using the OSIRIS Kn3 narrowband filter. A comparison between the narrowband and broadband spectra is shown in Figure 3, and both spectra agree within the 1$\sigma$ error bars. The Bolwer et al. spectrum decreases in flux more noticeably than the broadband spectrum; Bolwer et al. attribute this decrease in...
Figure 3. OSIRIS $H$- and $K$-band fluxes of HR8799b (scaled to 10 pc) plotted with $1\sigma$ uncertainties. The location of prominent water, CH$_4$, and CO absorption bands are indicated. The fluxes extracted without the speckle-suppression algorithm are shown as dotted lines. The bottom two curves are the mean residuals of fake planets with flat spectra extracted from the same data cubes (see the text for details). The Kn3 spectrum of Bowler et al. (2010) is shown as green pluses (scaled arbitrarily down) overplotted with the broadband spectrum (black points) interpolated onto the Bowler et al. Kn3 wavelength points. The mean $1\sigma$ uncertainties across the Kn3 range are shown at either end for each data set. The larger red filled symbols are the NICI CH$_4$ short/long (boxes), NIRC2 narrow (circles), and NIRC2 broadband (stars) photometry taken from Table 3. Blue symbols are the corresponding photometry derived from the OSIRIS spectra. (A color version of this figure is available in the online journal.)

Table 2

| $\lambda$ (\m$\mu$m) | $F_\nu$ (mJy) | Error (1$\sigma$) | $\lambda$ (\m$\mu$m) | $F_\nu$ (mJy) | Error (1$\sigma$) |
|----------------------|---------------|------------------|----------------------|---------------|------------------|
| 0.48                 | 0.25          | 0.08             | 1.97                 | 0.47          | 0.10             |
| 0.49                 | 0.44          | 0.04             | 1.99                 | 0.52          | 0.07             |
| 0.50                 | 0.42          | 0.06             | 2.00                 | 0.69          | 0.12             |
| 0.52                 | 0.36          | 0.03             | 2.02                 | 0.69          | 0.11             |
| 0.53                 | 0.33          | 0.08             | 2.04                 | 0.83          | 0.10             |
| 0.54                 | 0.47          | 0.05             | 2.05                 | 1.15          | 0.12             |
| 0.56                 | 0.53          | 0.08             | 2.07                 | 1.32          | 0.13             |
| 0.57                 | 0.57          | 0.06             | 2.08                 | 1.31          | 0.12             |
| 0.58                 | 0.63          | 0.06             | 2.10                 | 1.39          | 0.12             |
| 0.59                 | 0.73          | 0.04             | 2.12                 | 1.49          | 0.13             |
| 0.61                 | 0.76          | 0.04             | 2.13                 | 1.60          | 0.15             |
| 0.62                 | 0.83          | 0.05             | 2.15                 | 1.66          | 0.15             |
| 0.63                 | 0.89          | 0.03             | 2.16                 | 1.69          | 0.15             |
| 0.65                 | 0.85          | 0.05             | 2.18                 | 1.58          | 0.15             |
| 0.66                 | 0.85          | 0.07             | 2.20                 | 1.42          | 0.14             |
| 0.67                 | 0.90          | 0.04             | 2.21                 | 1.48          | 0.14             |
| 0.68                 | 1.05          | 0.04             | 2.23                 | 1.42          | 0.14             |
| 0.70                 | 1.02          | 0.06             | 2.24                 | 1.31          | 0.13             |
| 0.71                 | 1.00          | 0.04             | 2.26                 | 1.25          | 0.12             |
| 0.72                 | 0.89          | 0.04             | 2.28                 | 1.33          | 0.12             |
| 0.74                 | 0.74          | 0.04             | 2.29                 | 1.00          | 0.10             |
| 0.75                 | 0.60          | 0.06             | 2.31                 | 0.83          | 0.08             |
| 0.76                 | 0.58          | 0.05             | 2.32                 | 0.58          | 0.06             |
| 0.77                 | 0.43          | 0.05             | 2.34                 | 0.67          | 0.08             |
| 0.79                 | 0.37          | 0.05             | 2.36                 | 0.60          | 0.07             |

3. EMPIRICAL COMPARISONS

The $H$- and $K$-band spectra of HR8799b show several interesting features. In the $H$ band, a pronounced triangular shape indicative of weak collision-induced absorption (CIA) and low surface gravity is seen. The spectrum also shows no evidence of strong methane absorption (consistent with CH$_4$ on/off photometry) as would otherwise be expected in cold T-type BDs. At the $K$ band, the spectrum shows very deep water absorption bands. As in the $H$ band, there is no evidence of strong methane absorption or very strong CO absorption.

HR8799b occupies sparsely populated regions of near-IR color–magnitude diagrams (CMDs). Despite this fact, it is still useful to compare the observed spectra of HR8799b to those of known BDs. Figure 4 compares the $H$- and $K$-band spectra to a sequence of L and T dwarf spectra from type T6 to L1. The near-IR spectrum of HR8799b is fairly distinct from these typical field BDs, which generally have weaker water absorption and, for the T dwarfs, have deeper CH$_4$ absorption than seen in HR8799b. The fact that none of the hotter, L-type BDs provide a good match to HR8799b is consistent with the low effective temperature deduced from cooling track models (Marois et al. 2008).

BD near-IR colors exhibit considerable spread within a given spectral type and a number of peculiar (anomalously red or blue) L and T dwarfs are turning up (Kirkpatrick et al. 2008; Burningham et al. 2010). Consequently, while standard BDs do not match HR8799b very well, perhaps some of the more peculiar BDs might. The SpeX Prism spectral library provides over 300 low-resolution L and T dwarf spectra from $\sim$0.65 to

flux to possible weak CH$_4$ absorption. The broadband spectrum, which encompasses much more of the CH$_4$ absorption band, indicates an even weaker CH$_4$ signature than indicated by the Kn3 spectrum.

4 http://web.mit.edu/ajb/www/browndwarfs/spexprism
There is a clear correlation between the $J - K_s$ color and $\chi^2$. At first glance, this is simply restating that the best fits are found among the late L and early T dwarfs, which have the reddest $J - K_s$ colors. However, in this case, the best-fitting objects are noticeably redder than the mean $J - K_s$ color for their spectral type. The lower panel of Figure 5 shows $\chi^2$ versus $J - K_s$ color for all T dwarfs. For both bands, the best-fitting objects are noticeably redder than the mean $J - K_s$ color for their spectral type. While none of the objects available in the SpeX library are as red as HR8799b, there is a clear indication that early-T dwarfs, with anomalously red near-IR colors, are the best match. This trend was also independently noted by Bowler et al.

The best fit for the $H$ band, SDSS J151643.01+305344.4 (hereafter SDSS 1516+30), has $J - K_s = 1.77$ and has been classified as T0.5 $\pm$ 1 (Chiu et al. 2006; Burgasser et al. 2006) and T1.5 $\pm$ 2 (Burgasser et al. 2010). While SDSS 1516+30 matches HR8799b in the $H$ band very well, the comparison is poor at the $K$ band. Stephens et al. (2009), fitting their model atmospheres to available near-IR and mid-IR spectra of SDSS 1516+30, have also noted this object’s mid-IR colors are best reproduced by models with thick clouds and stronger CO absorption than predicted by pure chemical equilibrium, consistent with the high eddy diffusion coefficient used by Stephens et al. in their model comparison.

The best fit for the $K$ band is 2MASS J21392676+0220226 (hereafter 2M2139). The BD has $J - K_s = 1.68$ and has been classified as T0 (Reid et al. 2008), T1.5 (Burgasser et al. 2006), and recently as T2.5 $\pm$ 1 (Burgasser et al. 2010). Not only does 2.55 $\mu$m and includes a number of very red T dwarfs. These spectra were rebinned to match the resolution and sampling of the HR8799b OSIRIS spectra and used to find the best match to the $H$- and $K$-band spectra separately by minimizing $\chi^2$. All L and T dwarfs available in the SpeX library were used, except known binaries.

Figure 5 shows the results of fitting the SpeX observations to HR8799b. There is a clear $\chi^2$ minima between T0 and T2, with the best-fitting objects having spectral types T1.5 and T0.5 for the $H$ and $K$ bands, respectively (within the range found by Bowler et al., who fit their Kn3 spectrum to the same SpeX library). An empirical fit to the broadband photometry was also performed by Bowler et al.; their best overall match (a peculiar L6 2MASS 2148 + 4003) is compared to the $H$- and $K$-band spectra in Figure 5. While this L6 is a good match to the broadband photometry, it is not a good match to the broadband near-IR spectral data.

It is interesting to note that, while fitting BD spectra is completely independent of how well the broadband colors agree (each band was normalized individually for the $\chi^2$ calculation), there is a clear correlation between the $J - K_s$ color and $\chi^2$. At first glance, this is simply restating that the best fits are found among the late L and early T dwarfs, which have the reddest $J - K_s$ colors. However, in this case, the best-fitting objects are noticeably redder than the mean $J - K_s$ color for their spectral type. The lower panel of Figure 5 shows $\chi^2$ versus $J - K_s$ color for all T dwarfs. For both bands, the best-fitting objects (a T0.5 and a T1) are redder than the mean $J - K_s$ for T0 to T2 (vertical dotted lines, values taken from Faherty et al. 2009). While none of the objects available in the SpeX library are as red as HR8799b, there is a clear indication that early-T dwarfs, with anomalously red near-IR colors, are the best match. This trend was also independently noted by Bowler et al.

The best fit for the $H$ band, SDSS J151643.01+305344.4 (hereafter SDSS 1516+30), has $J - K_s = 1.77$ and has been classified as T0.5 $\pm$ 1 (Chiu et al. 2006; Burgasser et al. 2006) and T1.5 $\pm$ 2 (Burgasser et al. 2010). While SDSS 1516+30 matches HR8799b in the $H$ band very well, the comparison is poor at the $K$ band. Stephens et al. (2009), fitting their model atmospheres to available near-IR and mid-IR spectra of SDSS 1516+30, found $T_{\text{eff}} = 1000$–1100 K, log($g$) = 4.5 and a fairly cloudy atmosphere for such a low $T_{\text{eff}}$ ($f_{\text{red}} = 1–2$, in the Ackerman & Marley 2001 parlance). This $T_{\text{eff}}$ range is lower than typical for a T0.5 by 200–300 K (Stephens et al. 2009; Golimowski et al. 2004). Burgasser et al. (2010) identify SDSS 1516+30 as a weak, if not unlikely, binary candidate; however, binarity has yet to be tested with high-resolution imaging. Leggett et al. (2007) have also noted this object’s mid-IR colors are best reproduced by models with thick clouds and stronger CO absorption than predicted by pure chemical equilibrium, consistent with the high eddy diffusion coefficient used by Stephens et al. in their model comparison.

The best fit for the $K$ band is 2MASS J21392676+0220226 (hereafter 2M2139). The BD has $J - K_s = 1.68$ and has been classified as T0 (Reid et al. 2008), T1.5 (Burgasser et al. 2006), and recently as T2.5 $\pm$ 1 (Burgasser et al. 2010). Not only does
2M2139 match HR8799b reasonably well in the K band, the comparison is also about as good as SDSS 1516+30 in the H band. Burgasser et al. (2010) suggest that this object is a strong binary candidate, though no high-resolution imaging has been reported. It is quite possible that the object is single and simply non-standard with atmospheric properties that are a mix of those found in the cloudy and cloud-free BD photospheres. The K-band spectrum of 2M2139 is similar to that of SDSS J0758+42 (a T2), the BD found by Bowler et al. that best matched their Kn3 spectrum. The H-band spectrum of SDSS J0758+42, however, is noticeably bluer than the HR8799b OSIRIS H-band spectrum.

In summary, the empirical comparisons suggest a spectral type of T1 ±1, consistent with the T2 match that Bowler et al. (2010) found when comparing their Kn3 spectrum to the SpeX library. However, it appears as though only the reddest T dwarfs in this spectral type range provide reasonable matches to the HR8799b spectra. While a T1 is a close match, the near-IR colors of HR8799b are still well out of the normal range for this spectral type (being more consistent with the mid-L dwarfs; consequently, there are probably significant limitations to what any near-IR spectral type can tell us about the basic properties of HR8799b).

4. MODEL ATMOSPHERES

Recognizing that HR8799b falls outside of the color–magnitude range of most known substellar objects, comparing the observed data to synthetic spectra from model atmospheres is the next logical step. The PHOENIX model atmosphere code (Hauschildt et al. 1997) has been used to explore a variety of physical mechanisms that might explain the unusually red near-IR colors, low luminosity, and lack of strong methane bands. The version of PHOENIX (v16) used here has been substantially updated since the introduction of the PHOENIX “cond” and “dusty” BD atmosphere models (Allard et al. 2001). Some, but not all, of the updates include the replacement of the original Allard chemical equilibrium solver with a new, more robust and chemically complete, equation of state,6 a revised treatment of pure dust grains, improved alkali line profiles (Allard et al. 2007), and new molecular line lists, most notably for methane (Warmbier et al. 2009; Hauschildt et al. 2009) and water (Barber et al. 2006). Updates to the cloud modeling and the addition of mixing-induced disequilibrium chemistry are the most relevant here and are described in the following sections.

4.1. A Simple Cloud Model

Atmosphere models born from the study of BDs are commonly used to estimate the properties of giant planets and to estimate detection yields for direct imaging surveys, so it is natural to use such models as starting points for an analysis of

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6 The new equation of state routine is known as the Astrophysical Chemical Equilibrium Solver, or ACES (T. Barman & P. H. Hauschildt 2012, in preparation).
Recent improvements in the atmospheric modeling of transition BDs have been made, with most efforts attacking the cloud problem by adding one or more new parameters that regulate the cloud thickness, particle density, grain size distribution, and so forth (Ackerman & Marley 2001; Marley et al. 2002; Tsuji 2005; Burrows et al. 2006). However, see Helling et al. (2008a) for a completely different approach. The purpose of this paper is not to develop a new complex cloud model, but rather to identify the major physical processes that shape the spectral properties of HR8799b. To that end, a simple parameterized cloud model is adopted for the analysis presented here. The lower boundary (cloud base) is established by the intersection of the temperature–pressure (T–P) profile and the relevant chemical equilibrium condensation curve ($P_c$, using pressure as a proxy for height). The behavior of the cloud above $P_c$ is determined by a single free parameter ($P_{\text{min}}$). The equilibrium dust concentration is multiplied by a function that is 1 for $P_{\text{gas}} \geq P_{\text{min}}$ and decays exponentially for $P_{\text{gas}} < P_{\text{min}}$. If $P_{\text{min}} > P_c$, then the maximum dusty-to-gas ratio is also lowered relative to an equilibrium cloud. Examples of vertical cloud properties are discussed below in more detail. The particle sizes follow a log-normal distribution with a prescribed modal size ($a_0$) from 1 to 100 $\mu$m that is independent of height. This cloud model shares some similarities with other parameterized cloud models (Tsuji 2005; Burrows et al. 2006); however, rather than accounting for a small representative set of grains, all thermodynamically allowed grains with opacities in the PHOENIX databases are included in the total cloud opacity; see Ferguson et al. (2005) for a list of grains included in the model.

Figure 6 shows the location of HR8799b relative to L and T BDs in three near-IR CMDs. Synthetic color–magnitude sequences are also shown for chemical equilibrium models ($K_z = 0$) with radius set to that of Jupiter. Each sequence corresponds to a different $T_{\text{eff}}$ (with log($g$) = 4.0) and different cloud thickness varying from no clouds (at the blue end) to maximum cloud thickness (at the red end). These sequences of models are not associated with any evolutionary cooling tracks since the radius has been arbitrarily set to that of Jupiter. Such sequences illustrate the color trends produced by changing vertical cloud thickness (note this is different from the
uniform sinking of clouds). Figure 6 shows that the slope in a CMD across the L-to-T transition region, including the J-band brightening, is reproduced very well by atmospheres that have clouds smoothly decreasing in vertical thickness. However, as pointed out by Burrows et al. (2006), reproducing color trends does not guarantee good reproduction of near-IR spectra; the models shown in Figure 6 have not been compared to spectra of field BDs (an exercise left to another paper).

Figure 6 also shows that very red objects with relatively low luminosities, like HR8799b, can be produced by allowing low $T_{\text{eff}}$ models (typically assumed to be in the regime of cloud-free photospheres) to have clouds extending across their photosphere. The location of HR8799b in all three CMDs is intersected by a model with $T_{\text{eff}} = 1000$ K, log($g$) = 4.0, and a cloud that starts to fall off near 1 bar of pressure (though the precise $P_{\text{min}}$ matching the position of HR8799b is slightly different in each CMD). Figure 6 also suggests that the cloud properties of the HR8799 planets and 2M1207b are similar to those of late-type L dwarfs. The comparison in Figure 6 does not directly identify a best fit to the ensemble of data, but simply illustrates the impact that cloud thickness has on the precise properties of the HR8799 planets and 2M1207b are similar to those of late-type L dwarfs. The comparison in Figure 6 does not directly identify a best fit to the ensemble of data, but simply illustrates the impact that cloud thickness has on the near-IR colors of an object with typical substellar temperatures and gravities and points to a reasonable place to start looking for a good match to HR8799b, namely $T_{\text{eff}} \sim 1000$ K and log($g$) $\sim 4.0$.

4.2. Disequilibrium Chemistry and Vertical Mixing

The majority of atmosphere models for BDs and exoplanets assume local chemical equilibrium; in other words, the mole fractions for each chemical compound is determined by the local temperature and pressure and overall element abundances. For most atmospheric compounds, this is an excellent approximation as the chemical reactions that establish the equilibrium mole fractions are rapid compared to the mixing timescales (even in the convection zone). However, this is not always the case and it has been shown that, for example, the reactions that ultimately replace CO with CH$_4$ can be very slow in Jovian and BD photospheres. The same can be true for N$_2$ and NH$_3$. Prinn & Barshay (1977) showed that vertical mixing of CO can be important in Jupiter and, later, Fegley & Lodders (1996) predicted that this would also be the case for BDs. The spectroscopic detections of CO in several cool T-type BDs (Noll et al. 1997; Saumon et al. 2000, 2006; Geballe et al. 2009) along with weaker than predicted fluxes from ~4.5 $\mu$m ground and space-based photometry (Patten et al. 2006; Leggett et al. 2007) have established disequilibrium CO mole fractions as a common occurrence in BDs. Non-equilibrium chemistry is, therefore, a plausible mechanism for the apparently weak CH$_4$ absorption in HR8799b.

Disequilibrium CO/CH$_4$ and N$_2$/NH$_3$ chemistry has been added to PHOENIX following Smith (1998). The reaction timescales (in seconds) for N$_2$ $\rightarrow$ NH$_3$ and CO $\rightarrow$ CH$_4$ are from Lodders & Fegley (2002) and Yung et al. (1988), respectively:

$$\tau_{\text{N$_2$}} = \frac{1}{\kappa_{\text{N$_2$}}N(\text{H}_2)} \quad (1)$$

$$\kappa_{\text{N$_2$}} = 8.45 \times 10^{-8} \exp\left(\frac{-81515}{T}\right) \quad (2)$$

$$\tau_{\text{CO}} = \frac{N(\text{CO})}{\kappa_{\gamma}N(\text{H})N(\text{H}_2\text{CO})} \quad (3)$$

where $T$ is the gas temperature in Kelvin, $\kappa$ are the rate coefficients in cm$^3$ s$^{-1}$ (with $\kappa_{\gamma}$ interpolated from tabulated values in Yung et al. 1988), and $N$ (species) is the number density of the specified species in cm$^{-3}$.

The mixing timescale ($\tau_{\text{mix}}$) is computed using mixing-length theory in the convection zone, and set equal to $L_{\text{eff}}^2/K_{\text{L}}$ in the radiative zone, where the effective length scale ($L_{\text{eff}}$) is determined following the recipe in Smith (1998). For the range of models explored here, $L_{\text{eff}}$ was found to be $\sim 0.1-0.2$ $H_F$, with $H_F$ equal to the pressure scale height. This range of $L_{\text{eff}}$ is comparable to that found for Jupiter (Smith 1998; Visscher et al. 2010). The major unknown when modeling vertical mixing is the eddy diffusion coefficient ($K_{\text{L}}$) which, as is frequently done, is assumed here to be an adjustable parameter typically ranging from 100 to $10^8$ cm$^2$ s$^{-1}$. With these chemical and mixing timescales, the mole fractions of CO/CH$_4$ and N$_2$/NH$_3$ are “quenched” at a specific depth in the atmosphere where $\tau_{\text{mix}} = \tau_{\text{chem}}$. Below this point in the atmosphere, where $\tau_{\text{mix}} > \tau_{\text{chem}}$, chemical equilibrium prevails. This simple quenching model has been used in several recent studies of BDs and giant planets (Hubeny & Burrows 2007; Fortney et al. 2008).

The mixing and chemical timescales alone do not determine the final quenched mole fractions, the atmosphere temperatures and pressures still establish these values at the quenching depth. Consequently, effective temperature, surface gravity, and metallicity play an important role. Of particular importance is the relative location in the atmosphere of equal CO/CH$_4$ equilibrium mole fractions ($P_{\text{CO=CH$_4$}}$) and where $\tau_{\text{mix}} = \tau_{\text{chem}}$ ($P_q$). If $P_q > P_{\text{CO=CH$_4$}}$, the CO/CH$_4$ ratio can be greater than 1 in cool substellar atmospheres where CO would otherwise be predicted to be orders of magnitude less than CH$_4$.

As shown in Figure 7, the magnitude of the disequilibrium effect is sensitive to surface gravity, as this quantity acts to
shift the locations of $P_{g}$ and $P_{CO=CH_{4}}$. As gravity decreases, the crossing point of the radiative and mixing timescales moves to deeper layers (higher $T$), while the crossing of the equilibrium mole fractions of CO and CH$_{4}$ moves outward. For $K_{zz}$ values typically used in BD models and at low surface gravities expected for young exoplanets, the CO/CH$_{4}$ ratio can become significantly greater than 1—indicating a role reversal in the dominant carbon-bearing species at photospheric depths. Such a role reversal has not been explored thoroughly in BDs as their surface gravities and effective temperatures almost always place the timescale crossing point at a pressure lower than that of the equilibrium CO/CH$_{4}$ crossing point, resulting in a CO/CH$_{4}$ ratio always less than 1 (though still greater than equilibrium chemistry predictions). Hubeny & Burrows (2007) also noted a gravity dependence for non-equilibrium chemistry, but did not explore gravities below $\log(\right)$ and hotter ($T_{eff}$ = 7500 K) than the Sun and the UV flux from HR8799A may be sufficient to produce photochemical reactions in the atmosphere of HR8799b and certainly for the inner three planets. Based on a synthetic spectrum from a stellar model atmosphere calculation tailored for HR8799A (also using PHOENIX, with atoms and ions treated in non-local thermodynamic equilibrium), HR8799b is estimated to receive $\sim$8 times the flux that Jupiter receives from the Sun for $\lambda$ < 2000 Å and about 10 times as much flux as Neptune for $\lambda$ < 3000 Å. Given the larger atmospheric temperatures and possibly extreme non-equilibrium CO/CH$_{4}$ mole fractions compared to the giant planets in the solar system, the photochemical reactions in HR8799b will also be quite different from those found in the solar system. Consequently, photochemistry could impact the model comparisons discussed below. Unfortunately, photochemical reactions are not included in the version of PHOENIX used here. Future observations and modeling will be needed to address this possibility.

4.3. Photochemistry

HR8799b is located a mere 68 AU from its host star and receives far too little stellar flux for irradiation to have a significant influence on its atmospheric thermal structure (at near-IR photospheric depths). However, HR8799A is much more luminous ($\sim$5 $L_{\odot}$) and hotter ($T_{eff}$ = 7500 K) than the Sun and the UV flux from HR8799A may be sufficient to produce photochemical reactions in the atmosphere of HR8799b and certainly for the inner three planets. Based on a synthetic spectrum from a stellar model atmosphere calculation tailored for HR8799A (also using PHOENIX, with atoms and ions treated in non-local thermodynamic equilibrium), HR8799b is estimated to receive $\sim$8 times the flux that Jupiter receives from the Sun for $\lambda$ < 2000 Å and about 10 times as much flux as Neptune for $\lambda$ < 3000 Å. Given the larger atmospheric temperatures and possibly extreme non-equilibrium CO/CH$_{4}$ mole fractions compared to the giant planets in the solar system, the photochemical reactions in HR8799b will also be quite different from those found in the solar system. Consequently, photochemistry could impact the model comparisons discussed below. Unfortunately, photochemical reactions are not included in the version of PHOENIX used here. Future observations and modeling will be needed to address this possibility.

5. MODEL COMPARISONS

The complete set of HR8799b broadband photometry covers a broader wavelength range than the OSIRIS spectroscopic observations and, therefore, are likely to provide the greatest leverage when estimating effective temperature and, potentially when estimating the cloud properties. Marois et al. (2008) did not perform a model fit to the discovery photometry but, rather, focused on the parameters inferred from theoretical cooling tracks and the observed luminosity and age range. With new and improved photometry available, new model fits to these data are now warranted (see Bowler et al. 2010 and Currie et al. 2011 for independent model comparisons to similar photometric data and different model atmospheres). Additionally, the near-IR spectra are likely to contain the greatest information on surface gravity, chemical composition (via broadband absorption features), and also on the cloud properties. The goal in this section is to identify a single model atmosphere that best fits all three observational data sets and to understand what physical mechanisms play dominant roles in shaping the overall spectral energy distribution (SED).

5.1. Solar Abundance Model Comparison

The cloud and disequilibrium chemistry models described above were incorporated into PHOENIX. A modest grid of synthetic spectra was computed with five key parameters considered: $T_{eff}$ = [800 K–1500 K, 100 K increments], log($g$) = [2.5–5.0, in 0.5 increments and CGS units], $K_{zz}$ = $10^{4}$ cm$^{2}$ s$^{-1}$, $P_{min}$ = [$10^{5}$, $10^{6}$, 2 $\times$ $10^{6}$, 4 $\times$ $10^{6}$, 6 $\times$ $10^{6}$, 8 $\times$ $10^{6}$, 1 $\times$ $10^{7}$ dynes cm$^{-2}$], and $\delta_{0}$ = [1, 5, 10, 100] $\mu$m. Solar abundances were adopted for this grid (Asplund et al. 2005), but the question of non-solar metallicity will be discussed below.

Given the large number of free parameters and the sparse sampling of the model grid, obtaining a traditional fit by $\chi^{2}$ minimization that samples a continuous distribution of parameters is beyond the scope of this paper. Instead, a $\chi^{2}$ was computed for each model spectrum from the grid compared to each of the three data sets (photometry, $H$- and $K$-band spectra). When computing a $\chi^{2}$ for the photometric comparison, each model spectrum was convolved with the appropriate filter response functions to produce synthetic photometry in $J$, $H$, $K_{s}$, and $L'$ bands along with the filters used by Hinz et al. (2010) and Currie et al. (2011); $M_{\text{bar}}, z$, and a filter centered at 3.3 $\mu$m. When comparing models to the OSIRIS spectra, the synthetic spectra were first smoothed (with a Gaussian kernel) to the native instrumental response for OSIRIS ($R$ $\sim$ 4000), then binned down to the same wavelength sampling (and in the same manner) as done for the OSIRIS data cubes. The best-fitting model was taken to be the model with smallest mean $\chi^{2}$ for the three data sets. Figure 8 shows the mean reduced-$\chi^{2}$ distribution across $T_{eff}$ and log($g$) for the model comparisons to all three sets of data. The best fit is located at $T_{eff}$ = 1100 K and log($g$) = 3.5. Formal error bars for the fit have not been determined; however, models within the mean $\chi^{2}$ = 4 contour are all reasonable fits to the data. Therefore, the error on the fit is close to $\pm$100 K for effective temperature and $\pm$0.5 dec for surface gravity and comparable to most BD model fit uncertainties found in the literature.

Figure 9 compares the best fit to each of the three data sets. The model spectrum agrees very well with the photometry. The differences seen near $J$-band slopes, $H$ and $K_{s}$, and $L'$ are within the $1\sigma$ error bars. The model only slightly underpredicts the flux at 3.3 $\mu$m, compared to the Currie et al. (2011) detection. The flux measurement at 3.3 $\mu$m excludes a large number of models, especially at high $T_{eff}$, and is quite sensitive to the presence or absence of CH$_{4}$. Observations of other systems (e.g., 2M1207b) in this bandpass would be very useful.

At the $H$ band, the model spectrum is marginally consistent with the observed spectrum. The slope at the red and blue ends are not as steep as observed and, overall, the model spectrum is too flat. It is possible that the residual flux from the bright speckle feature seen near $\sim$1.7 $\mu$m (see Figure 3) remains in the final observed spectrum and is responsible for the larger disagreement seen at this wavelength. As for the red and blue $H$-band slopes, these wavelength regions have historically been difficult to fit mainly due to missing CH$_{4}$ line opacity. Consequently, the level of disagreement at the $H$ band is not too surprising. At the $K$ band, the model compares favorably to the observed spectrum. The red and blue slopes are in good agreement as is the location of the peak flux.

When the data sets are compared to the models individually, slightly different results are obtained. The best-fitting model
to the photometry has $T_{\text{eff}} = 1000$ K and $\log(g) = 5.0$. The $H$ band is best fit by a model with $T_{\text{eff}} = 1200$ K and $\log(g) = 3.0$ (clearly favoring low gravity); however, this model has colors that are too blue when compared to observed photometry, especially in the $z$ band. At the $K$ band, the best fit has $T_{\text{eff}} = 1100$ K and $\log(g) = 3.0$. When the mean $\chi^2$ for $H$ and $K$ is minimized, the best fit is $T_{\text{eff}} = 1200$ K and $\log(g) = 3.5$. Each of these best-fit models has a slightly different cloud thickness, but all have $K_{zz} = 10^4$ cm$^2$ s$^{-1}$ and $a_0 = 5$ $\mu$m.

5.2. Inferred Atmospheric Properties

Despite the imperfect fit to the $H$ band, there is strong evidence for low surface gravity. In Figure 10, the $H$-band OSIRIS spectrum is compared to a sequence of model spectra with different gravities but with all other parameters equal to those of the overall best fit shown in Figure 9. The observed triangular-shaped pseudo-continuum is best reproduced by the lowest-gravity model. The shape of the $H$ band is often associated with T dwarfs. This gravity indicator is decreased for field T dwarfs as CH$_4$ opacity dominates the near-IR and competes with that of H$_2$O around the $H$ band. However, if CH$_4$ mole fractions are reduced by vertical mixing and non-equilibrium chemistry at low gravity and low $T_{\text{eff}}$, as is likely the case for HR8799b, then the behavior of the $H$ band with gravity seen in hotter L dwarfs will persist even down to effective temperatures most often associated with T dwarfs.

Water is by far the main molecular opacity source shaping the OSIRIS spectra. The combination of high $K_{zz}$ and low $\log(g)$ leads to, in the models used here, a sharp reduction in the CH$_4$ mixing ratio (e.g., see Figure 7) with a correspondingly large increase in the CO mixing ratio. Figure 11 compares the best-fit model to those with the same parameters and temperature–pressure structure, but with CO or CH$_4$ opacities removed. Weak CH$_4$ absorption is evident in the $K$ band, as is a possible contribution from CO; however, the strongest evidence for CH$_4$ absorption is the Currie et al. (2011) detection (and Hinz et al. 2010 upper limit) at 3.3 $\mu$m where the flux is clearly suppressed by CH$_4$. Unfortunately, the Hinz et al. (2010) upper limit near 4.5 $\mu$m is not sensitive enough to distinguish between the majority of the models explored here and, therefore, does not provide strong constraints on CO.

The best-fitting, solar abundance, atmosphere model for HR8799b (shown in Figure 9) has a cloud base just below $P =$
Barman et al.

Figure 9. Comparison between the best-fitting solar abundance model and the broadband photometry (top panel) and OSIRIS broadband H and K spectra (middle and bottom panels). Observations are shown in black with $1\sigma$ error bars, the model spectrum in red and model photometry in green. The photometric upper limit from Hinz et al. (2010) is plotted as a horizontal bar. (A color version of this figure is available in the online journal.)

1 bar with a composition that is typical for BD models with $T_{\text{eff}} > 1000$ K. Magnesium-silicate (e.g., forsterite) and iron grains are the most abundant, with a mixture of other grains with lower mole fractions. The structure of the cloud has many characteristics in common with BD cloud predictions, namely an abrupt drop in the gas-to-dust ratio at higher pressures and temperatures, where condensation is no longer thermodynamically favored. This cloud base is determined by the intersection of the $T$–$P$ profile and the relevant condensation curves (see Figure 12). At lower temperatures and pressures, the gas-to-dust ratio drops exponentially (by design in this study, see lower panel of Figure 12) in a similar manner as the smooth decrease in dust predicted by the models of Helling & Woitke and Marley & Ackerman (see Helling et al. 2008b for a comparison). Location and vertical extent are the most important aspects of the cloud structure for HR8799b preferred by this work. Generally, the cloud base retreats to larger pressures with decreasing $T_{\text{eff}}$ and fixed gravity; however, the low gravity of the best-fitting HR8799b model counters this trend. As gravity decreases, the $T$–$P$ profile shifts to lower pressure, moving the cloud base along with it. At $\log(g) \sim 3.5$, the cloud base sits near the photosphere with vertical extent that is far less than an equilibrium cloud model (see Figure 12), but sufficient to produce redder colors than cloud-free BD models with the same $T_{\text{eff}}$ and $\log(g)$.

Figure 10. Model gravity sequence compared to the OSIRIS H-band spectrum. All model parameters are identical the best-fitting model shown in Figure 9 except for gravity, which varies from $\log(g) = 3.5$–5.0 (cgs units). The triangular shape of the H band is best explained by low surface gravity. (A color version of this figure is available in the online journal.)

The modal particle size for HR8799b, favored by this work is around 5–10 $\mu$m. Grains of this size are within the range predicted by the Marley & Ackerman and Helling & Woitke for
models with $T_{\text{eff}} \sim 1000$ K. Both groups predict sizes between 1 and 100 μm at high pressures (>10 bar) and sizes that decrease below a micron at low pressures (<1 bar). Burrows et al. (2006) chose $a_0 = 100$ μm (at the high end of most predictions) for their set of reference model atmospheres arguing that, in general, L-type BDs are best fit when larger grain sizes are assumed but acknowledged that such large grains make fitting the reddest late L dwarfs difficult. As already discussed above, the observed near-IR spectra of the reddest late L and early T dwarfs come closest to matching the $H$- and $K$-band spectra of HR8799b and, thus, perhaps it should not be surprising that the best model fits are found among those with smaller grains. However, the simulations of Cooper et al. (2003) predict that particle size should actually increase with decreasing gravity. When comparing grain sizes adopted by different models, it is important to remember that grain size alone does not determine the total cloud opacity—composition also plays an important role. In the cloud model used here, a mixture of grains are included in the total opacity and such a cloud can be more opaque than a cloud including opacity from only pure silicate grains of the same size.

6. DISCREPANCY BETWEEN ATMOSPHERE AND EVOLUTIONARY MODELS

Since the distance to HR8799 has been precisely measured (39.4 ± 1.0 pc; van Leeuwen 2007), the luminosity of HR8799b is known ($\log L_{\text{obs}}/L_{\odot} = -5.1$) and, therefore, radii can be assigned to each $T_{\text{eff}}$ value on the left vertical axis of Figure 8; these radii are given by the right, nonlinear, vertical axis. Using these radii and the $\log(g)$ axis values, lines of constant mass can be plotted that are completely independent of interior cooling track models. Dynamical stability arguments suggest that, in order for the three planets to be as massive as indicated by the cooling tracks, they must be in double 2:1 resonances and have masses $\leq 20 M_{\text{Jup}}$ (Fabrycky & Murray-Clay 2010). Using $20 M_{\text{Jup}}$ as an upper limit on the mass of HR8799b sets a $T_{\text{eff}}/\log(g)$ boundary in Figure 8 (rightmost solid line). This limit was the motivation for not extending the model grid to $\log(g)$ greater than 5.0. A second boundary, at lower surface gravity, can be set using the Kelvin–Helmholtz timescale ($\tau_{\text{KH}}$), which reflects the approximate scaling of detailed planetary cooling curves. Since the planets almost certainly are older than 1 Myr, $\tau_{\text{KH}} = 1$ Myr is a very conservative limit and is indicated by the leftmost dotted line in Figure 8. In other words, planets located to the left of this dotted line could not maintain the observed luminosity of HR8799b for more than 1 Myr (under the usual assumption that the primary energy source is gravitational potential energy). This timescale limit was used to exclude a second minimum at $\log(g) = 2.5$ near $T_{\text{eff}} = 1100–1200$ K. For $T_{\text{eff}} = 1100$ K ± 100 K, the radius must be 0.75 $R_{\text{Jup}}+0.17$ to match $L_{\text{obs}}$. Using this radius and the best-fit surface gravity, $\log(g) = 3.5 ± 0.5$, the corresponding mass is just under 1 Jupiter mass ($\sim 0.72 M_{\text{Jup}}-0.6$), a factor
of \( \sim 10 \) lower than predicted by evolutionary cooling tracks. A radius of 0.75 \( R_{\text{Jup}} \) is also \( \sim 30\% \) smaller than cooling track predictions.

While the implied mass of \( \sim 0.7 \, M_{\text{Jup}} \) is comfortably within the ranged needed for dynamical stability, it is difficult to form objects, either by core accretion or by gravitational instability, with this mass, radius (\( \sim 0.75 \, R_{\text{Jup}} \)), and the planet’s observed luminosity. This combination is inconsistent with our understanding of degenerate or partially degenerate H+He gaseous bodies with any reasonable heavy-metal composition or of any age (Baraffe et al. 2008). Figure 13 compares the \( T_{\text{eff}} \) and \( \log(g) \) values from the atmosphere model fit to those predicted by evolution models from Baraffe et al. (2003). While the best-fit surface gravity is within the expected range, an effective temperature of 1100 K is \( \sim 300 \) K above the temperatures predicted by all current hot-star cooling tracks, for an age less than 100 Myr. The high effective temperature would be consistent with evolution predictions if the system was old (>1 Gyr) and \( \log(g) \) was greater than \( \sim 5 \). However, as shown above, such a high surface gravity would be very difficult to reconcile with the \( H \)-band spectrum and would have serious dynamical stability implications (Fabrycky & Murray-Clay 2010; Marois et al. 2010; Currie et al. 2011).

Giant planet evolution models are known to become increasingly unreliable at young ages for the simple reason that planet formation is not well understood and, thus, the initial conditions are not known. This fact was highlighted by Baraffe et al. (2002) with a focus on the formation of BDs and low-mass stars from the collapse of molecular clouds (the main realm of applicability for the common cond and dusty cooling tracks). More recently, Marley et al. (2007) revisited the question of initial conditions by comparing the evolution of planets starting with high \( T_{\text{eff}} \) and luminosity ("hot start," as often adopted for BDs and motivated by the Stahler birth line) to the evolution of planets starting with the colder and less luminous initial conditions ("cold start") that emerge from core-accretion planet formation models. Marley et al. concluded that young planets that formed by core accretion are cooler, smaller, and thus fainter than planets formed by single-mode gravitational collapse within a disk or giant molecular cloud. Given that HR8799b is clearly underluminous compared to the hot-start models and overluminous for the cold-start models (Marley et al. 2007; Fortney et al. 2008), it is unlikely that initial conditions alone are responsible for the offset between the atmosphere-inferred properties and the evolution track predictions (a situation similar to that of 2M1207b, as discussed below).

A \( \sim 300 \) K offset in effective temperature is similar in magnitude to offsets in effective temperature found for a number of M-, L- and T-type binary BDs (Liu et al. 2008, 2010; Konopacky et al. 2010; Dupuy et al. 2010). This suggests that a systematic offset may exist between the atmosphere and interior model predictions. A 10%–20% difference in \( T_{\text{eff}} \) is comparable to the differences between cloudy and cloud-free tracks in the 10 and 100 Myr range and 10 \( M_{\text{Jup}} \) (Chabrier et al. 2000), so it

![Figure 12](image_url)
is still unclear what fraction of such an offset could be resolved by improvements purely in the modeling of the interior or the atmosphere. However, the increasing number of free parameters in substellar atmosphere models is a clear indication that they are even more complex than previously imagined. Thus, the possibility still remains that hidden among the free parameters is an atmosphere model with $T_{\text{eff}}$ and $\log(g)$ that matches the evolution model predictions.

### 6.1. Non-solar Abundance Model Comparisons

The need for an intermediate cloud model (constrained vertically) to match the observed broadband photometry of all three HR8799 planets was first discussed by Marois et al. (2008). Marois et al. (2008) showed that a synthetic spectrum from a solar abundance model atmosphere with an intermediately thick cloud and $T_{\text{eff}}/\log(g)$ consistent with evolution model predictions ($820 \, \text{K}/4.0$) matched the available observed photometry at $3\sigma$. While only CH4, J, H, Ks, and L' photometry were available in 2008, this earlier (chemical equilibrium) model remains consistent ($\sim 3\sigma$) with all of the revised/new broadband photometry (see Figure 14). However, the true level of agreement between this model and reality is revealed by the OSIRIS data (Figure 14) where much weaker CH4 absorption is observed than predicted by this chemical equilibrium model. This discrepancy highlights the difficulties in obtaining a reliable description of young giant planets from photometry alone, which can hide significant clues about their atmospheric properties.

To explore differences between the atmosphere and evolution derived bulk parameter more closely, a separate grid of atmosphere models was computed, with $T_{\text{eff}}$ and $\log(g)$ uniformly sampling the evolution track predictions that are consistent with the observed luminosity of HR8799b (down the center of the gray band in Figure 13). The clouds and non-equilibrium chemistry were modeled in the same manner as in the solar abundance grid described above, yet only $a_0 = 5$ and 10 $\mu$m were used. The metallicity, however, was extended beyond solar abundances to include two metal-rich cases ([Fe/H] = 0.5 and 1.0) and an abundance pattern matching the observed peculiar ($\lambda$-boo-type) subsolar abundances of HR8799 (Sadakane 2006). The best-fitting model from this “evolution-consistent” grid has $T_{\text{eff}} = 896$ K, $\log(g) = 4.3$, [Fe/H] = 1.0, $\log L_{\text{bol}}/L_{\odot} = -5.06$ (see Figure 14). The agreement between this metal-rich model and the OSIRIS $H$- and $K$-band spectra is similar to that of the previous hotter, solar abundance, best fit. The broadband photometric comparison is also good (mostly to $1\sigma$), but slightly worse compared to the solar abundance model at $J$ and $K$ bands (see Figure 14) but still consistent at $2\sigma$. The lowering of $T_{\text{eff}}$ by increasing metallicity is to be expected. Looking at the $M_J$ versus $J-K$ CMD in Figure 6, in order for cooler atmospheres to match HR8799b they need to be redder in $J-K$ and this can be accomplished by increasing metallicity and/or by decreasing the modal grain size (Burrows et al. 2006). The mean $\chi^2$ for this metal-rich comparison is $\sim 5.5$ and only marginally better than the solar abundance model at the same temperature and gravity. However, given that even in this smaller grid the
parameter space is still sparsely sampled, there is plenty of room for improvement.

6.2. Comparison to 2M1207b

In near-IR CMDs, the object most similar to HR8799b is 2M1207b, a planet–mass companion to a young (\~{}10 Myr) BD (Chauvin et al. 2005). In addition to their very red colors, these two objects also share the same puzzling fact that their SEDs appear hotter than evolution models predict for their ages and luminosities. Available photometry and spectroscopy of 2M1207b show all the signs of an object with effective temperature \~{}1600 K (Mohanty et al. 2007; Ducourant et al. 2008; Patience et al. 2010). When combined with current luminosity estimates (\log(L/L_\odot) \sim{} -4.5), the radius implied by this high effective temperature is \~{}0.7 \text{R}_{\text{Jup}}, comparable to the small size inferred above for HR8799b. Consequently, both 2M1207b and HR8799b are underluminous compared to hot-start evolution models and overluminous for cold-start models, for their respective ages.

Mohanty et al. (2007) concluded that the high effective temperature was likely correct and, thus, 2M1207b is underluminous by \~{}2.5 mag. Given that the distance to the source (52.4 \pm 1.1 pc) is known (Ducourant et al. 2008), Mohanty et al. (2007) proposed that an edge-on disk surrounds 2M1207b providing the necessary extinction. Alternatively, Mamajek & Meyer (2007) proposed that 2M1207b is not a normal gas giant but rather the product of a recent protoplanetary collision and the observed flux is the afterglow from the hot, physically small, remnant. Given that HR8799b suffers from a similar quandary, it is unlikely that such rare events should be necessary to explain the first and only two objects found in the very red and faint regions of near-IR CMDs.

Given the puzzling similarities between 2M1207b and HR8799b mentioned above, it is useful to compare their SEDs. Moderate resolution (R \sim{} 500–1500) H- and K-band spectra of 2M1207b (Patience et al. 2010) are compared to the HR8799b spectra in Figure 15. The spectrum of 2M1207b has shallower H$_2$O bands with a distinct CO bandhead at 2.3 \text{\mu}m and no prominent CH$_4$ absorption—a spectrum very much like typical L dwarfs. The only feature the two sets of spectra have in common is a triangular-shaped spectrum in the H band, consistent with both objects having low surface gravity (indicative of low mass and/or youth). Assuming both objects are not the result of recent protoplanetary collisions, 2M1207b and HR8799b may well represent important evolutionary states of substellar atmospheres. Independent of the edge-on disk hypothesis, if the ages of the two systems are correct, then the planets’ masses are probably the same to within a factor of two. If the masses are similar, then these objects provide evidence for rapid spectral changes from 10 to \sim{} 30 Myr for young planet-mass objects.
Figure 15. OSIRIS $H$- and $K$-band spectra of HR8799b compared to the spectra of 2M1207b (Patience et al. 2010) binned to the same resolution and interpolated onto the same wavelength grid as the OSIRIS observations. The 2M1207b spectra were independently normalized for this comparison. The error bars for the 2M1207b spectra are significantly smaller than those of HR8799b and, for clarity, are not shown. (A color version of this figure is available in the online journal.)

The analysis of HR8799b described above provides a third possible explanation for the properties of 2M1207b—namely, that certain combinations of cloud coverage and non-equilibrium chemistry (and perhaps metallicity) are capable of producing a very low temperature object with an L-type near-IR spectrum. The two hypotheses proposed by Mohanty et al. and Mamajek et al. hinge on the assumption that the effective temperature of 2M1207b is $\sim 1600$ K. However, looking back at Figure 6, 2M1207b is intersected by an intermediate cloudy model with $T_{\text{eff}} = 1100$ K, close to evolution model predictions and consistent with a typical radius. Also, given the apparently extreme reduction of CH$_4$ in the atmosphere of HR8799b, it is conceivable that the CO/CH$_4$ mole fractions are quenched at even higher pressures in the younger (and likely lower gravity) atmosphere of 2M1207b, resulting in a low temperature atmosphere that is rich in CO (perhaps enough to produce a distinct CO bandhead at 2.3 $\mu$m, as seen in the spectrum from Patience et al.) and without significant amounts of CH$_4$. A search for such a cold, cloudy, model atmosphere with extreme non-equilibrium chemistry for 2M1207b is underway and will be presented in a separate paper (T. Barman et al. 2012, in preparation).

Despite their spectral differences, the most straightforward explanation for the positions of 2M1207b and HR8799b in near-IR CMDs is one of atmospheric origins. Instead of an edge-on disk or recent protoplanetary collision, a full exploration of the cloud properties, non-equilibrium chemistry, and metallicity is likely required to reproduce their observed SEDs with a model atmosphere having $T_{\text{eff}}$ and $\log(g)$ consistent with evolution predictions. However, as more and more important physical processes are identified in substellar atmospheres, the number of free (and essentially independent) parameters increases and the ability to exhaustively explore all possible parameter combinations diminishes—more work (along the lines of Freytag et al. 2010) is clearly required to reduce the current set of parameters and produce more unified model atmospheres.

7. CONCLUSIONS

The HR8799 planetary system provides a rare opportunity to study multiple coeval giant planets at very young ages. The data and model comparisons described above are only first steps toward understanding just one of these planets and many challenging problems remain. While the data presented here offer the broadest spectral coverage for an HR8799 planet obtained so far, the low resolution and modest signal-to-noise set strong limits on what can be understood about HR8799b. In addition to the limits imposed by the data, the complexity and excess of free parameters in model atmospheres for giant planets and BDs also set strong limits on what can be confidently inferred from observations. With these limitations in mind, several conclusions can be drawn from the observations of HR8799b.

The SED of HR8799b is shaped by water absorption bands that are deeper than typically seen in late L and early T field dwarfs, but clearly consistent with a hydrogen-rich planetary atmosphere. Only weak methane absorption is detected in the $K$ band along with very weak CO absorption, indicating strong non-equilibrium chemistry and vertical mixing. Additional evidence for very weak methane absorption is seen in the $H$-band spectrum (and narrowband photometry). The near-IR spectra and broadband photometry are smoothed and reddened by photospheric clouds not typical for BDs with $T_{\text{eff}} \leq 1100$ K. The triangular shape of the $H$-band spectrum clearly indicates low surface gravity ($\log(g) < 5$), consistent with an age significantly younger than that proposed by Moya et al. (2010).

The best match from a solar abundance grid of model atmospheres (having $T_{\text{eff}} = 1100$ K and $\log(g) = 3.5$) found above implies a radius that is too small to be consistent with both the observed luminosity and the current understanding of the interiors and evolution of H+He-rich planets, regardless of mass, age, or interior core-size. Furthermore, a variety of model atmospheres can be constructed that fit the observed broadband photometry very well (at least from the model grids.
used here), yet do not match the observed spectra — indicating that parameters inferred from broadband photometric fits alone should be treated with caution.

While certainly not perfect, the model comparisons presented here are better representations of the observations than those presented by Bowler et al. (2010) and Currie et al. (2011). In particular, Bowler et al. found fits to their narrowband spectrum that require $T_{\text{eff}} > 1100$ K, which leads to unphysically small radii of $R < 0.75 R_{\odot}$. These previous studies also had difficulty simultaneously fitting the photometry; for example, Currie et al. found it difficult to match most of the photometry, especially at $\lambda > 3 \mu$m (even with their thick or patchy cloud models) and Bowler et al. had difficulty matching the near-IR colors for reasonable $T_{\text{eff}}$ values. These difficulties may stem from the sparse sampling of model parameters in the grids used in these two studies (e.g., those parameters regulating non-equilibrium chemistry and clouds), though considerable degeneracies are present when fitting only broadband photometry. Furthermore, the best-fitting models for HR8799b found by Currie et al. typically have deeper methane absorption than seen in the broad $K$-band spectrum described above or the narrowband spectrum of Bowler et al. Interestingly, Currie et al. emphasize cloud thickness/coverage over non-equilibrium chemistry, in particular as it pertains to the near-IR SED. The analysis presented here, however, requires deep quenching of CO/CH$_4$ to match the near-IR spectra (and near-IR narrowband photometry), placing equal importance on both atmospheric effects when reproducing the observations.

Finding a solar abundance atmosphere that matches the complete set of data and interior/cooling track predictions seems unlikely (though still not inconceivable). A more likely scenario is that the atmosphere of HR8799b has an elemental composition different from the Sun. A model atmosphere that matches the observations reasonably well (as well or better than any previous study), including the predictions of evolution models, can be found if the atmosphere is enhanced by a factor 10 in metals compared to the Sun. Such a model has $T_{\text{eff}} = 896$ K and $\log(g) = 4.3$, consistent with a young giant planet with mass less than $10 M_{\oplus}$. At $\sim 30$ Myr, as recently proposed by Zuckerman et al. (2011), the mass of HR8799b would be $4$ to $5 M_{\oplus}$.

If HR8799b is indeed metal-rich and near the very low end of the original mass range inferred by Marois et al. (2008), then core accretion may be the preferred formation mechanism for the HR8799 planetary system. Metal enrichment, relative to the host star, is often considered a hallmark of the core-accretion scenario and recent work indicates that if the HR8799 planets formed by gravitation instabilities, they should show little metal enhancement (Helled & Bodenheimer 2010). Atmosphere models with either solar abundances or the metal-poor $\lambda$-boötis-like abundances of the star result in $T_{\text{eff}}$ that are too high and corresponding radii that are too small. The analysis of Bowler et al. (2010) also indicates that enhanced metals (and low gravity) are likely to produce the best model fits.

While these conclusions rest heavily on models not tested by a large population of young giant planets (such a population should emerge in the coming years), the conclusions are self-supporting. Low surface gravity, clearly preferred by the model fitting (regardless of metallicity) and indicated by the shape of the $H$-band spectrum, may well allow for more efficient vertical mixing. As surface gravity decreases, the radiative–convection boundary moves closer to the photosphere, the maximum convective velocity increases, and $K_{\text{eff}}$ also increases, at least in the convection zone (Freytag et al. 2010; Rice et al. 2011).

Enhanced vertical mixing leads to enhanced non-equilibrium chemistry, as suggested by the absence of CH$_4$ absorption in both $H$ and $K$ bands. However, the connection between gravity and mixing is merely suggestive as the primary physical mechanism for mixing into the radiative zone unknown. Low surface gravity may also establish atmospheric conditions (not found in the higher gravity field T dwarfs) that support thick clouds at photospheric depths, allowing the intersection of the $T$–$P$ curve and the condensation curve to be near the photosphere even at low $T_{\text{eff}}$. A connection between clouds and surface gravity is already starting to emerge for BDs. Evidence for a gravity dependence for the L-to-T transition has already been reported (Metchev & Hillenbrand 2006; Saumon & Marley 2008) and, recently, Stephens et al. (2009) found that low-gravity BDs in their sample remain cloudy at lower effective temperatures, with the L-to-T transition temperature decreasing from 1300 K to 1100 K as $\log(g)$ decreases from 5.0 to 4.5. Metchev & Hillenbrand (2006) also found that HD20303B, a low-gravity transition BD, is $\sim 250$ K cooler than field objects with similar spectral types. If HR8799b has a $\log(g) \sim 4$, then it is possible that this emerging trend extends to this object as well, and would suggest that very low gravity objects (like young giant planets) are cloudy down to temperatures well below 1000 K. The structure of the clouds needed to reproduce the colors and spectra of the HR8799 planets and 2M1207b are similar to those needed for late L dwarfs. This suggests that such objects, from an atmospheric point of view, are extensions of the L-type field dwarfs rather than members of a new class altogether. This conclusion is distinct from that of Currie et al. who claim the clouds must be significantly thicker than those required to match L dwarfs. Finally, enhanced metals would also lead to enhanced grain formation and conditions ripe for atmospheric cloud formation. The combination of these atmospheric conditions lead to bulk properties, inferred from the SED analysis alone, that include a radius consistent with the observed luminosity and current understanding of giant planet interiors as well as a mass that fits comfortably within the limits required for long-term dynamical stability.

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