Head-Tail Galaxies: Beacons of High-Density Regions in Clusters

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ABSTRACT
Using radio data at 1.4 GHz from the ATCA we identify five head-tail (HT) galaxies in the central region of the Horologium-Reticulum Supercluster (HRS). Physical parameters of the HT galaxies were determined along with substructure in the HRS to probe the relationship between environment and radio properties. Using a density enhancement technique applied to 582 spectroscopic measurements in the $2^\circ \times 2^\circ$ region about A3125/A3128, we find all five HT galaxies reside in regions of extremely high density ($>100$ galaxies/Mpc$^3$). In fact, the environments surrounding HT galaxies are statistically denser than those environments surrounding non-HT galaxies and among the densest environments in a cluster. Additionally, the HT galaxies are found in regions of enhanced X-ray emission and we show that the enhanced density continues out to substructure groups of 10 members. We propose that it is the high densities that allow ram pressure to bend the HT galaxies as opposed to previously proposed mechanisms relying on exceptionally high peculiar velocities.

Key words: radio continuum; galaxies; galaxies: clusters: general; galaxies: clusters: individual: A3125, A3128

1 INTRODUCTION
The term “radio galaxy” is usually used to describe a class of objects which have characteristic jets of highly relativistic particles blasted out from the host galaxy. Although most if not all galaxies have some radio emission, the term “radio galaxy” is usually reserved for those galaxies that have a high radio-to-optical luminosity ratio (Kellermann et al. 1989).

Head-tail (HT) galaxies are a sub-class of radio galaxy whose radio jets are distorted such that they appear to bend in a common direction. This gives rise to the so-called “head-tail” morphology where the radio jets are bent back to resemble a tail with the luminous galaxy as the “head”. (The term “head-tail” was coined by Miley et al. (1972) but the first HT galaxy was discovered by Ryle & Windram (1968).) Although considered somewhat exotic in their early years, HT galaxies are now considered quite common.

HT galaxies are generally FRI in morphology (Venturi et al. 1998) which, assuming that the FRI/FRII break is caused by environment, is consistent with HT galaxies being found in clusters as their denser environments effectively restrict the outflow of the jets (Owen & Ledlow 1994).

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The bent morphology of jets in HT galaxies has always been attributed to one of two possible mechanisms relating the relative velocity of the host galaxy to the intra-cluster medium (ICM). In the first mechanism the host galaxy has a larger than expected peculiar velocity, leading to ram pressure on the jets (Miley et al. 1972; Rudnick & Owen 1976). Where ram pressure, \(P_{\text{ram}}\), is defined as:
where $\rho$ is the density of the ICM, and $v$ is the relative velocity between the galaxy and the ICM (Gunn & Gott 1972).

In the second mechanism, strong winds in the ICM, believed to be caused by dynamical interactions such as cluster-cluster mergers or accretion of groups on to clusters, bend the jets. The phenomenon is often referred to as “cluster weather” (Burns 1998). These mechanisms have been invoked to explain the difference in the degree to which the jets are bent, with Narrow-Angle Tailed galaxies (NATs) believed to be formed via ram pressure in cases where the host galaxy has a peculiar velocity of the order of 600 $\text{km} \, \text{s}^{-1}$ and is embedded in a hot, dense plasma (Venkatesan et al. 1994), and Wide-Angle Tailed galaxies (WATs) believed to be the result of “cluster weather” (Klamer et al. 2004). NATs are most likely to be caused by the passage of the galaxy through the cluster (Odea & Owen 1985) while WATs are more likely to be associated with dominant cluster galaxies (Owen & Rudnick 1979).

Although the distinction between ram pressure and ICM effects, such as “cluster weather” (Burns 1998), is often drawn in literature, we note that cluster weather still implies ram pressure as the cause of the jet bending. The true distinction between the ram pressure theory and the cluster weather theory lies in the velocities of the HT galaxies relative to the kinematic centre. Henceforth, when we refer to the ram pressure theory we are indicating that the galaxy has a high peculiar velocity and is in fact travelling at a velocity which differs significantly from the statistical distribution of velocities for other galaxies in the cluster. Cluster weather on the other hand will mean that the galaxy does not have a significantly different velocity, but rather it is the motion of the ICM that is causing the bent morphology.

In both cases it is the relative motion of the host galaxy with respect to the ICM that causes the bending. Consequently when only viewing the radio morphology it is not possible to distinguish between a galaxy with high peculiar velocity moving through an undynamical, stationary ICM, and a galaxy with the ICM rushing past. However, spectroscopic redshifts can be used to determine the distribution of velocities for cluster galaxies and, via a statistical analysis, the relative importance of the host galaxy’s motion with respect to its neighbours can be established.

This paper explores the environment of a sample of HT galaxies and examines their usefulness as a tool to investigate dynamical conditions in the host clusters. We examine the well-studied major merger A3125/A3128 which resides in the central region of the Horologium-Reticulum Supercluster (HRS). Using a large spectroscopic dataset in conjunction with multi-frequency radio imaging we present a quantitative analysis of the environment of HT galaxies as compared to samples of both other cluster galaxies and other cluster radio galaxies. The paper is laid out as follows: Section 2 discusses the spectroscopic and radio data in the region surrounding A3125/A3128 and the parameters associated with the detected HT galaxies; Section 3 presents the substructure detection; Section 4 gives the results of our statistical analysis of the environments surrounding HT galaxies and Sections 5 and 6 present the discussion and conclusions.

This paper uses $H_0 = 70 \, \text{km} \, \text{s}^{-1} \, \text{Mpc}^{-1}$ and $\Omega_M = 0.27$. This gives a scale of 1.191 kpc/arcsecond at 18500 $\text{km} \, \text{s}^{-1}$, the mean redshift of the HRS.

## 2 HEAD-TAIL GALAXIES: DETECTION AND MORPHOLOGY

To investigate the relationship between HT galaxies and the dynamics of the cluster, we study HT galaxies in A3125/A3128, a major merger in the central region of the Horologium-Reticulum supercluster (HRS). The HRS is the second largest supercluster in the local universe (out to 300 Mpc), preceded only by the Shapley supercluster. The cores of rich superclusters are ideal environments to study cluster mergers because the high peculiar velocities induced by the enhanced local density favours cluster-cluster and cluster-group collisions (Venturi et al. 2004).

### 2.1 Spectroscopic Data

We compiled a catalogue of 2181 spectroscopic redshifts in the central region of the HRS. Data were obtained from all previously published surveys (Rose et al. 2002; Fleenor 2006) in addition to an unpublished catalogue of spectroscopic data obtained at ANU by T. Mathams in 1989 and with data from the NASA Extragalactic Database (NED). A spatially and velocity limited subset of the 2181 galaxies covering $2^\circ \times 2^\circ$ in area and 15,000 km s$^{-1}$ < $cz$ < 22,000 km s$^{-1}$ in velocity centred on the major merger A3125/A3128 yields 582 redshifts.

When compiling a catalogue from several different sources which have overlapping observations of the same galaxies, care must be taken to ensure that individual objects do not appear more than once in the final catalogue or that close, but independent objects are not incorrectly conflated. To avoid these problems we employ a visual inspection method for all cases where objects from different sources appear within the spatial and recessional uncertainties of the catalogues ($\theta'' \sim 400$ km s$^{-1}$). In such cases postage stamps from the Digitized Sky Survey (DSS) are inspected. This is an important step as a completely automated compilation method might inadvertently remove true close-proximity galaxies.

The final catalogue of 582 galaxies has a magnitude limit of $m_b=23$ and is $\sim 90\%$ complete to $m_b=18$ and $\sim 60\%$ complete to $m_b=19.5$. The optical coverage of the HRS region appears uniform (see Figure 1).

The resultant spectroscopic catalogue, when combined with the available radio images, provides an ideal dataset with which to investigate the environ-
Figure 1. Top left: Skyplot of the HRS region with $m_b$ cutoff = 18. Top right: Skyplot of the HRS region with $m_b$ cutoff = 18.5. Bottom left: Skyplot of the HRS region with $m_b$ cutoff = 19. Bottom right: Skyplot of the HRS region with $m_b$ cutoff = 19.5. Contours are set at 10% intervals. The crosses mark the positions of the HT galaxies.

Figure 2. Histogram of galaxy velocities from our position and velocity limited catalogue (of 582 galaxies). The shaded histogram plots the 46 radio galaxies. The velocities of our HT galaxies are shown with dotted lines. It can be seen that the HT galaxies are clearly in the main component of A3125/A3128, as are the radio galaxies.

2.2 Radio Observations

Radio images at 1.4 GHz of the central 3.28 square degrees surrounding the A3125/A3128 system (centred at 03:29:45, -53:00:00) were obtained with the Australia Telescope Compact Array (ATCA). The observation was a 40-pointing, linear mosaic consisted of $2 \times 128$ MHz observations centred at 1344 and 1432 MHz. The data were obtained in three different array configurations over 39 hours in 2001 (see Table 1). The primary calibrator was PKS B1934-638 and...
the phase calibrator was PKS B0302-623. The synthesized beam size is 9.90 by 7.87 arcseconds with a position angle of 31.9°. The mean RMS noise for the image (Figure 3) is 0.1 mJy beam$^{-1}$ and lowest contours are set at 0.3 mJy beam$^{-1}$, three times the RMS noise. Over 500 radio galaxies are detected in the 1.4 GHz data (Johnston-Hollitt et al. 2008) of which 46 are identified with optical galaxies having redshifts found in our spectroscopic catalogue that placed them in A3125/A3128.

Low frequency data for the region were also available from the Sydney University Molonglo Sky Survey (SUMSS) (Mauch et al. 2003), which used the Molonglo Observatory Synthesis Telescope (MOST) to assemble a comprehensive picture of the Southern radio sky ($\delta \leq -30°$) at 843 MHz. Data obtained with the MOST have a 3 MHz bandwidth and the resultant images from SUMSS have a resolution of the order of 1'.

2.3 Identifying HT Galaxies

Five HT galaxies are identified in the 1.4 GHz data, including two that were previously detected (Johnston-Hollitt et al. 2004). Typically only one or two HT galaxies tend to be found in non-relaxed clusters in the current literature. However, the available radio surveys covering clusters tend either to be quite shallow, or have low sensitivity to extended emission (e.g. the NVSS (Condon et al. 1998)) and hence only the brighter extended sources are detected. In a comparably shallow survey to those commonly used we would only be able to see the two HT galaxies previously studied by Johnston-Hollitt et al. (2004) as they are the brightest HT galaxies of our sample. Interestingly, Millon (2003) surveyed 7 square degrees in the central region of the Shapley Supercluster to a similar depth to our data (80 $\mu$Jy beam$^{-1}$ at a resolution of 16''), and also found five HT galaxies.

The optical and radio parameters of the HT galaxies are discussed below and Table 2 summarises the respective frequency data.

2.3.1 Optical Counterparts

The positional data for the HT galaxies is taken as the coordinates of their optical counterparts; these values were obtained from NED. The HT galaxies all lie at comparable distances and have similar magnitudes. The velocities of the HT galaxies are marked by dotted lines in Figure 2.

2.3.2 Flux Densities

We used the Karma (Gooch 1996) package Kvis to determine the flux densities of the HT galaxies at both available frequencies. The error in the flux density is

$$ \Delta S = \sqrt{(Sb)^2 + a^2}, $$

where $S$ is the flux density, $b$ is the residual calibration error, and $a$ is the RMS noise.

Because the flux density of our radio galaxies is so high relative to the RMS noise, the uncertainty in total flux density is dominated by the 10% residual calibration error typical for the ATCA.

2.3.3 Power and Fanaroff-Riley Classification

Using the radio power formula from Hogg (1999) we calculate the radio powers for the five HT galaxies. Their powers range between $8.37 \times 10^{22}$ W Hz$^{-1}$ and $4.46 \times 10^{24}$ W Hz$^{-1}$. We find all of our HT galaxies to be either FRI or at the transition between FRI and FRII, which is consistent given the absolute magnitudes of these galaxies (Owen et al. 2000). This is in agreement with previous literature (e.g. Venturi et al. 1998).

3 SUBSTRUCTURE DETECTION

Substructure, which may be defined as the presence of two or more clumps of matter within a cluster, is a clear sign of incomplete relaxation and is hence indicative of the dynamical nature of a cluster of galaxies (Pinkney et al. 1994). Substructure can be detected in two-dimensions via analysis of X-ray surface brightness data, but to investigate the full three-dimensional nature of a cluster a large spectroscopic survey is required.

By identifying substructure in the cluster we can investigate the relationship between the location of the HT galaxies and the underlying dynamical structure. Caldwell & Rose (1993) showed that the A3125/A3128 system is in a post-merger state and that A3125, the smaller of the two clusters, is more dispersed due to its recent passage through A3128.

Substructure in the central region of the HRS has been studied previously by Rose et al. (2002). Using their spectroscopic data for 532 galaxies in this region, Rose et al. (2002) identified four groups on the high velocity side of A3125/A3128. These groups were selected on the basis of their clumping in either RA-cz space or Dec-cz space. We note that in this context “groups” refer to substructure groups and not galaxy groups. Furthermore, Rose et al. (2002) noted the presence of a void 4000 km s$^{-1}$ on either side of the clusters. They interpret this as a consequence of the large gravitational potential of the cluster and hence the voids are indicative of galaxies being absorbed from the peripheries into the clusters.

The substructure detection method used by Rose et al. (2002) is limited in that groups can only be oriented along the RA, Dec or cz axes. This means any “falling groups” that are falling “diagonally into the cluster”, for example, will not be as easily detectable.

We believe that rather than identifying substructure in the A3125/A3128 region, Rose et al. (2002) have merely identified parts of the cluster. For this reason we have chosen not to repeat the method used by Rose et al (2002) when detecting substructure. Instead we use density enhancements in our catalogue.
Table 1. Details of the 1.4 GHz ATCA observations. Columns 1 and 2 give the telescope configuration and observing date respectively. Column 3 is the integration time in minutes, column 4 is bandwidth in MHz, column 5 gives the central frequencies for each IF and column 6 is the uv-range covered in kilowavelengths. Observations were carried out by R. W. Hunstead, M. Johnston-Hollitt, J. A. Rose, W. Christiansen, & M. Fleenor. Full details of the observations and reduction are available in [Johnston-Hollitt et al. 2008].

| Configuration | Date                  | Integration (min) | Bandwidth (MHz) | Frequencies (MHz) | UV-range (kλ) |
|---------------|-----------------------|-------------------|-----------------|-------------------|--------------|
| 6A            | 8th & 10th December, 2001 | 880               | $2 \times 128$ | 1344 & 1432       | 0.50 - 20.72 |
| 6D            | 27th November, 2001    | 840               | $2 \times 128$ | 1344 & 1432       | 1.57 - 27.72 |
| 1.5D          | 19th & 20th November, 2001 | 640             | $2 \times 128$ | 1344 & 1432       | 0.36 - 27.43 |

Figure 3. Central region of the 3.28 square degrees 1.4 GHz image of A3125/A3128 overlaid on optical data from the DSS. The RMS of the image is 0.1 mJy beam$^{-1}$ and the contour is at three times the RMS noise. The five HT galaxies identified in the image are denoted with boxes. We note that Galaxy A appears to have a companion galaxy nestled between its lobes. This galaxy is in fact in the background and hence not considered in this study. The X-ray peaks of A3125 (lower right with three peaks) and A3128 (upper left with two peaks) are marked with crosses - the size indicates the strength of the X-ray emission.
to determine the location of substructure in this region. Using a physical parameter such as density enhancement provides a more accurate analysis of the substructure within A3128/A3125.

3.1 3D Density Enhancements

To identify the location of substructure in A3125/A3128 we introduce our own method of identifying substructure using density enhancements.

We searched for density enhancements in our catalogue of galaxies using the HIPASS group-finder, a group finding program written by Stevens (2005). This group finder needs only sky position and velocity to operate, and finds groups using the hierarchical algorithm (Tully 1987).

Specifically, the group-finder uses the grouping equations of Gourgoulhon, Chamaraux, & Fouqu\`e (1992) (hereafter GCF92) to determine the likely three-dimensional separation of two galaxies using their on-sky positions and recessional velocities. The GCF92 equations compute separation in three regimes:

(i) When the velocities of the two galaxies differ by less than some user-specified parameter $V_1$, the velocity difference is considered to be entirely peculiar, and thus the separation is calculated using only plane-of-sky information.

(ii) When the velocities of the two galaxies differ by more than some user-specified parameter $V_0$, then the galaxies are considered to be ungroupable; i.e. they are separated by Hubble expansion and are not likely to be bound to each other.

(iii) When the velocities of the two galaxies differ by an amount between these two extremes, then the separation is calculated using a combination of plane-of-sky and velocity information.

At the start of the procedure, each galaxy is considered as an individual “unit”. The group-finder calculates the separation between every pair of units using the GCF92 equations, and from this the radius of the groups (the average galaxy-galaxy separation of the group members) that would result from each pair.

The galaxy number density for each potential group is then calculated simply as the total number of galaxies in the two units divided by the volume. This differs from the method of GCF92, who calculated mass density, but this is the only significant difference.

The pair of units with the highest number density are then grouped into a single unit, which replaces the two original units. The separations between each of the remaining pairs of units in the catalogue are then recalculated and the process begins again. It should be noted that a unit with more than one galaxy does not hide any information: when the radius of a group is calculated, the group-finder always considers the individual galaxies. The group-finder stops when the number density of the densest pair is below some user-specified stopping density. For each group that it identifies, the group-finder calculates both the radius as calculated from the GCF92 equations and the projected radius, using just the plane-of-sky information.

Because the group-finder uses number density to identify groups, it can be thought of as a type of nearest-neighbour finder. Indeed, if $V_1$ is set to a sufficiently large value, and $V_0$ was set to the same value, the group-finder would use only plane-of-sky information and would identify only those groups which appear as dense regions on the sky. However setting $V_1$ lower will allow the group-finder to identify structure along the line of sight.

A comparison could be made to the non-parametric algorithm DEDICA (Pisani 1993, 1996), which was designed to identify clusters and determine whether they have substructure. However, the work presented in this paper is less concerned with the identification of the clusters themselves, and more concerned with the environment immediately surrounding each galaxy in the cluster. As the hierarchical algorithm gives information about clustering strength at all scales, it is particularly suited for our purposes.

3.2 2D Density Enhancements

To minimize the uncertainties that accumulate when using velocity as a pseudo-spatial parameter, we repeat the density enhancement study using projected radii instead of 3D density. That is, the substructure groups are still selected using exactly the same method as above, but we measure the projected radii of the galaxies.
Figure 4. Results of the HIPASS group finder at $V_l = 1000$ km s$^{-1}$. The x-axis gives the density in galaxies/Mpc$^3$ on a logarithmic scale while the y-axis is the number of galaxies that can be found in groups.

groups to determine relatively whether HT galaxies reside in different environments to non-HT galaxies.

For each galaxy in the catalogue, as more galaxies are grouped with it, we can track how the projected radii increase as well, in effect showing us how density changes with increasing group membership.

4 RESULTS

From Figure 2 we can deduce that all the HT galaxies lie in the main peak of the velocity distribution. Thus it appears that the HT morphology is not a consequence of high peculiar velocity.

Figure 4 shows the fraction of the spectroscopic catalogue assigned to a group as a function of stopping density. It is clear that at very low densities ($\sim 10^{-2}$ galaxies/Mpc$^3$), all the galaxies in the catalogue are placed in groups and at very high densities ($\sim 10^5$ galaxies/Mpc$^3$) almost no galaxies are in groups as expected.

The cut-off densities of the HT galaxies – that is, the maximum density of the groups containing a HT galaxy – show that HT galaxies reside in extremely dense regions. Table 3 gives the cut-off densities for each of the HT galaxies. We note that due to incomplete optical coverage of our region not all redshifts have been determined hence the cut-off densities are lower limits. This is especially true for Galaxies A and C where there clearly exists close galaxies, but without reliable redshifts these galaxies are not incorporated in our cut-off density calculation.

In fact, we find that the HT galaxies are in the densest regions of the cluster. We assume that the density of the environment has no bearing on whether a galaxy is radio-emitting (as per Venturi et al. (2000)).

4.1 Density Cut-Offs

To ensure that the extremely dense environments surrounding our HT galaxies are statistically significant, we perform Monte Carlo shuffles on our density cut-offs.

Monte Carlo shuffles work by calculating the mean (or median) density cut-off of the five HT galaxies. We then choose five random galaxies and average their density cut-offs. This is repeated 100000 times and the fraction of shuffles that produces density cut-offs greater than the mean (or median) density cut-off of the HT galaxies show the statistical significance of the HT galaxies’ density cut-offs.

If HT galaxies are truly in denser regions than non-HT galaxies we would expect that

- HT galaxies are found in statistically denser regions than non-HT galaxies with radio emission,
- HT galaxies may or may not be found in statistically denser regions than non-HT galaxies without radio emission. This is because galaxies without radio emission can potentially be in denser regions than HT galaxies but as they have no radio morphology, a comparison is meaningless, and
- radio galaxies are not found in statistically different regions to non-radio galaxies (Venturi et al. 2000).

Consequently, we perform Monte Carlo shuffles on three different sets of sample populations.

(i) HT v non-HT radio
(ii) HT v non-HT
(iii) radio v non-radio

The results are shown in Table 4. A result is considered statistically significant when the Monte Carlo fraction is less than or equal to 0.01.

The values generated for HT v non-HT radio and HT v non-HT are statistically significant for both average and median statistics which means that HT galaxies do reside in regions of enhanced density. Using the mean, it appears that there is a significant difference between radio and non-radio galaxies, but this

| Test                  | Mean | Median |
|-----------------------|------|--------|
| HT v non-HT radio     | <0.001 | 0.001 |
| HT v non-HT           | <0.001 | <0.001 |
| non-radio v radio     | 0.011 | 0.081 |

Table 3. Cut-off densities for HT galaxies.

| HT galaxy | $V_l = 1000$ km s$^{-1}$ |
|-----------|--------------------------|
| A         | 645                       |
| B         | 60680                     |
| C         | 135                       |
| D         | 29072                     |
| E         | 64919                     |

Table 4. Fraction of Monte Carlo Shuffles that have cut-off densities greater than or equal to the mean (or median) HT cut-off density.
Table 5. Radii of the groups containing HT galaxies. Column 1 is the HT galaxy and the remaining columns give the projected radius of the group (in kpc) with the specified population.

| HT galaxy | Group population |
|-----------|-----------------|
|           | 2   | 3   | 4   | 5   | 10  | 15  | 20  | 30  |
| A         | 71  | 126 | 174 | 223 | 402 | 540 | 848 | 1092|
| B         | 16  | 29  | 43  | 93  | 264 | 357 | 558 | 860 |
| C         | 119 | 139 | 158 | 177 | 273 | 360 | 474 | 1172|
| D         | 23  | 53  | 105 | 134 | 392 | 591 | 721 | 1222|
| E         | 15  | 81  | 108 | 141 | 296 | 517 | 811 | 1127|

Table 6. Average radii for all groups. Column 1 is the class of galaxy (HT: the five HT galaxies; radio: the 46 radio galaxies (including the five HT galaxies); all), and the remaining columns give the average radii of the groups (in kpc) with the specified population.

| Class | Group population |
|-------|-----------------|
|       | 2   | 3   | 4   | 5   | 10  | 15  | 20  | 30  |
| HT    | 49  | 86  | 118 | 154 | 325 | 471 | 682 | 1095|
| Radio | 95  | 153 | 211 | 261 | 504 | 674 | 751 | 1055|
| All   | 109 | 163 | 216 | 268 | 489 | 645 | 749 | 1020|

significance disappears when using the median statistic. We attribute this to the very high densities of the HT galaxies, which could skew the radio mean, but would have little effect on the median.

4.2 Radii

To verify the results from the density parameter, and to explore the environment further out from each of the HT galaxies, we conduct an analysis using projected radius. For this, the groups are still identified using the HIPASS group finder in the usual way, however we now examine their projected radii as the number of galaxies in the group (group population) increases.

Table 5 gives the projected radii of the groups found around each HT galaxy, while Table 6 shows the average projected radii of the groups found around all the galaxies in the catalogue, and is broken up according to galaxy class (ie. HT, radio, all). The average group radii for HT galaxy groups is roughly half that of those around other classes of galaxies, for groups of less than 10 members. This suggests that HT galaxies are surrounded by closer companions than non-HT galaxies, and supports the result that HT galaxies reside in denser regions of the cluster.

We perform Monte Carlo shuffles on our data to determine whether the smaller radii associated with HT galaxy groups is statistically significant for groups with 10 or fewer galaxies. There does not appear to be a significant difference between the groups around radio galaxies and those around all galaxies. The slightly different values at the group population of two may be attributed to the smaller number of radio galaxies, and thus the likelihood of randomly selecting a HT galaxy group is higher.

Table 7. The fraction of Monte Carlo Shuffles that produced five groups with an average radius less than the average radius of the HT galaxy groups, when all groups are considered and when only groups around radio galaxies are considered.

| Group Population | All | Radio |
|------------------|-----|-------|
| 2                | 0.004 | 0.045 |
| 3                | 0.007 | 0.010 |
| 4                | 0.005 | 0.009 |
| 5                | 0.008 | 0.010 |
| 10               | 0.021 | 0.023 |
| 15               | 0.073 | 0.053 |
| 20               | 0.244 | 0.160 |
| 30               | 0.891 | 0.735 |

From the Monte Carlo Shuffles we can see that the smaller radii associated with HT galaxy groups is statistically significant for groups with 10 or fewer galaxies. There does not appear to be a significant difference between the groups around radio galaxies and those around all galaxies. The slightly different values at the group population of two may be attributed to the smaller number of radio galaxies, and thus the likelihood of randomly selecting a HT galaxy group is higher.

5 DISCUSSION

5.1 Close Companions to the HT Galaxies

Looking more closely at the high density groupings, we find that each HT galaxy has a close companion. Using DSS images, we locate the closest galaxy as identified by the HIPASS group finder; these companion galaxies are shown on radio-optical overlays in Figure 5. Each HT galaxy has a close companion with projected separations ranging from 15 to 119 kpc.

5.2 Analysis of Statistics

In our Results section we showed that the densities surrounding HT galaxies were significantly higher than those surrounding non-HT galaxies. In the following sections we will discuss the implications of this result.

It has been well documented in the literature that there is no correlation between the density of a region and the number of radio galaxies observed (e.g. Venturi et al. 2000). Consequently the marginally significant P-values obtained for the non-radio v radio mean density test are startling as it suggests that radio galaxies exist preferentially in dense environments!

We are reluctant to believe that radio galaxies exist preferentially in dense environments. If anything, we expect the densities of the non-HT galaxies to be higher than the non-HT radio galaxies as high density non-radio galaxies could potentially skew our data. We
Figure 5. The HT galaxies and their companion galaxies. The HT galaxy is shown in the “zoomed out” image while the optical galaxy and it’s companion are shown in the “zoomed in” image. The 1.4 GHz ATCA data is red while the 843 MHz MOST data (Mauch et al. 2003) is blue. For both frequencies the lowest contour is three times the image RMS noise (0.1 mJy beam$^{-1}$ at 1.4 GHz) and contours increase at $\sqrt{2}$ intervals. As noted earlier, the radio galaxy nestled between the lobes of Galaxy A is a background source and not considered in this study. The green square identifies the closest companion to the HT galaxies as identified by eye. We note that this is the same as the companion identified using the HIPASS group finder in three cases. We do not have redshifts for the companion galaxies to Galaxies A and C - this is a possible explanation for their lower cut-off densities. The arrow indicates the hypothesized trajectory of the galaxy.
attribute this bias to the difference in optical and radio coverage. The optical data have an average coverage of $\sim90\%$ down to a magnitude of 18. Conversely, our radio coverage is complete to flux densities of $\geq0.5$ mJy at 1.4 GHz. This is essentially $100\%$ complete for HT galaxies as the faintest HT galaxy has a flux density of 11 mJy at 1.4 GHz. The incomplete optical catalogue has the consequence of skewing the density of the optical catalogue downwards; that is, it appears less dense.

From Figure 1 we can see that the optical coverage is $100\%$ down to $m_b=18$ over most of the field and the average coverage decreases uniformly over the field with fainter magnitude cutoffs. This means that the optical coverage surrounding the HT galaxies is not any different to the optical coverage surrounding the other radio galaxies. Thus, we think that the HT v non-HT test is free of optical biases.

The p-values obtained for both HT galaxy tests strongly suggest that the densities surrounding HT galaxies are significant and hence we are confident (to $\sim99\%$) that HT galaxies exist in significantly denser environments than non-HT galaxies.

The results derived from the 2D density enhancement analysis using projected radii appear to support the results we obtain from our 3D density enhancement study.

Our results also show that the tighter-knit environments surrounding HT galaxies remain significant for group populations of up to 10 galaxies. This suggests that high densities are required for HT galaxy morphology to occur.

5.3 A Brief Foray into the Shapley Supercluster

To ensure the density values obtained for the HT galaxies in the central region of the HRS are realistic, we test our process on the closest supercluster in the local universe: the Shapley supercluster. We do not have data for this region so we rely on the radio data available from Miller (2003) and Venturi et al. (2000). From their papers we identify five HT galaxies in the three central clusters of the Shapley supercluster. (The three clusters are a well-studied merger A3556-A3558-A3562, not dissimilar to A3125/A3128 in the HRS.) The positional and redshift catalogue is taken straight from NED, so our visual catalogue compilation method is not utilized. The optical coverage of the Shapley supercluster is comparable to our data for the HRS (Venturi et al. 2000). Furthermore, the sensitivity of the radio data is also comparable to our radio data so we will assume the two superclusters may be directly compared. The Shapley supercluster has roughly double the galaxy density of the HRS (Pronst et al. 2003).

The density cut-off values for the HT galaxies in the Shapley supercluster are similar to the values obtained for the HT galaxies in the HRS (Table 8).

The values for the Shapley supercluster are slightly higher than for the HRS but this may be attributed to the denser environment of the Shapley supercluster. However, the similarity of the results for a peculiar velocity of 1000 km s$^{-1}$ in both the HRS and the Shapley supercluster gives us confidence in the velocity of our results. Thus we may say with confidence that HT galaxies do form in regions of enhanced density.

### Table 8. Cut-off densities for HT galaxies in Shapley.

| HT galaxy      | Cut-off density (galaxies/Mpc$^3$) |
|----------------|-----------------------------------|
| J132357-313845 | 8639                              |
| J133331-314058 | 1708                              |
| J132802-314521 | 183654                            |
| J133542-315354 | 367                               |
| J133048-314325 | 16729                             |

5.4 Further Evidence for High-Density Environments

Inspection of archival ROSAT X-ray data shows that four of the five HT galaxies (A to D) reside in regions of enhanced emission near the centres of A3125 and A3128. There is also evidence that the fifth galaxy (galaxy E) is in a region of X-ray emission which is above the X-ray background but it is at a much lower level than for the other galaxies. Since X-ray emission also traces regions of enhanced density these results support the statistical analysis presented above. Interestingly, Burns et al. (1994) found a high correlation between X-ray emission and radio galaxies, and that X-ray emission around WATs was generally more extended. Furthermore, Miller & Owen (1999) found that X-ray emission around FRI galaxies is more extended than that around FRIs. Our results are consistent with these findings. Details of the X-ray emission will be discussed in a separate paper.

The morphology of the HT galaxies is itself indicative of a dense environment, as it is believed that the jets of FRI galaxies are less efficient in transporting energy due to the dense ICM surrounding them (Fanaroff & Riley 1974).

The morphology of HT galaxies has always been believed to be due to either:

- the host galaxy having a high peculiar velocity (radio trail theory, Miley et al., 1972), or
- the host galaxy being swept by strong winds (cluster weather, Burns, 1998).

Using the optical data for our HT galaxies we find that the velocities are not exceptionally different to the cluster mean and hence it would appear that the host galaxies are victims of cluster weather. However, studies of cluster weather indicate that if cluster winds were the sole cause of the HT morphology, the velocity required for ram pressure is less than previously thought (Pinkney, 1995). While we find that the velocities of the HT galaxies are not especially great,
the densities in which they reside are among the densest regions of the cluster. Although we measure the number density of galaxies, this value follows from the cluster potential, and thus should also be representative of the density of the ICM due to gravitational attraction.

We propose that the morphology is due to ram pressure effects, but it is the exceptionally high gas densities that contribute to the morphology, as opposed to peculiar velocities or high velocity winds due to cluster weather (as $P_{\text{ram}} \propto \rho v^2$).

6 CONCLUSIONS

We have compiled the largest positional and spectral catalogue of the major merger A3125/A3128. This well studied merger lies in the central region of the HRS, the second largest supercluster in the local universe. Using 1.4 GHz radio data we identified five HT galaxies in this region.

To investigate the relationship between the HT galaxies and their underlying dynamical structure we identified substructure in the central region of the HRS. We found, using the HIPASS group finder, that the HT galaxies all exist in some of the densest regions of the cluster merger system A3125/A3128. We performed a series of statistical tests to probe three populations of galaxies within A3125/A3128 (15,000 kms$^{-1}$ $\leq$ cz $\leq$ 35,000 kms$^{-1}$); all galaxies (582), all radio galaxies with a spectroscopic redshift within A3125/A3128 (46) and all HT galaxies (5). Results of the single-tailed permutation tests showed, to 99% confidence, that HT galaxies reside in significantly denser environments than either non-HT galaxies or radio galaxies which are known members of the clusters. These results are also replicated in a comparative study of the five HT galaxies in the centre of the Shapley Supercluster.

Additionally, we found each of the HT galaxies in A3125/A3128 reside in regions of enhanced X-ray emission and have a very close companion (ranging from 15 - 119 kpc in projected separation). As the peculiar velocities of the host galaxies are not significantly different to the mean cluster velocity this suggests that it is the density of the environment which plays a significant part in the bending of the jets. However, the presence of the close companions is also likely to play a role in at least the motion of the host galaxies, though it is likely that the morphology of the jets on kpc-scales is more closely related to ram pressure and hence cluster weather. Thus, we propose that the combination of both a close companion and the cluster environment contribute to the bent morphology of HT galaxies.

Ultimately, we have found that HT galaxies are excellent indicators of dense regions and dynamical interactions within galaxy clusters and thus are clear signposts of “cluster weather”. As a result, HT galaxies may be used as barometers for the storms that occur in the cluster environment, although due to the complication of close companion galaxies, reading these barometers may not be as easy as first thought. We propose that the diffuse tails, on at least kpc-scales, can be used as giant wind socks and are effective beacons of cluster weather. Detection of the full extent of HT galaxies requires low frequency observations of reasonable resolution. Instruments such as the GMRT and LOFAR will provide large samples of HT galaxies which may then be used as beacons of high-density regions in clusters.

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