We present the transmission spectra of the hot Jupiter HD 209458b taken with the Space Telescope Imaging Spectrograph aboard the Hubble Space Telescope. Our analysis combines data at two resolutions and applies a complete pixel-by-pixel limb-darkening correction to fully reveal the spectral line shapes of atmospheric absorption features. Terrestrial-based Na I and H I contamination are identified that mask the strong exoplanetary absorption signature in the Na core, which we find reaches total absorption levels of $\approx 0.11\%$ in a 4.4 Å band. The Na spectral line profile is characterized by a wide absorption profile at the lowest absorption depths and a sharp transition to a narrow absorption profile at higher absorption values. The transmission spectra also show the presence of an additional absorber at $\approx 6250$ Å, observed at both medium and low resolutions. We performed various limb-darkening tests, including using high-precision limb-darkening measurements of the Sun to characterize a general trend of ATLAS models to slightly overestimate the amount of limb darkening at all wavelengths, likely due to the limitations of the model’s one-dimensional nature. We conclude that, despite these limitations, ATLAS models can still successfully model limb darkening in high signal-to-noise ratio transits of solar-type stars, like HD 209458, to a high level of precision over the entire optical regime (3000–10000 Å) at transit phases between second and third contact.

Subject headings: methods: data analysis — planetary systems — stars: individual (HD 209458)
position of the star within the slit. The shifts are typically on the order of ~0.1 pixel within an orbit (the typical guiding accuracy of HST) and approximately ±0.5 pixel from orbit to orbit. Several iterations of measuring pixel shifts, applying those shifts, and producing a new template were performed in order to create a reliable final template spectra for cross-correlation. With this method, we were able to align the spectra to ~0.1 pixel. With different Doppler and pixel shifts for each spectrum, the PSF of the stellar absorption line profiles can be reconstructed with a much greater sampling than can be obtained from an individual spectrum (see Fig. 1).

The extracted STIS spectra were then used to create a photometric time series covering the G750M bandpass by integrating the flux from each exposure. The resulting photometric light curve exhibits all the systematic instrumental effects noted in Brown (2001). As in both Brown (2001) and Charbonneau et al. (2002), we corrected for these systematic effects by discarding the first exposure of each orbit, discarding the first orbit of each visit, and applying flux corrections by removing a fourth-order polynomial fit to the photometric time series phased on the HST orbital period. The photometric corrections were then applied to the spectra themselves, such that different photometric bands could easily be produced.

The Na D wavelength region is also covered by HST STIS G750L observations, described by Knutson et al. (2007), although at much lower resolution. The G750L data set was also reduced, in the same manner as the G750M. The cross-correlation analysis, however, was only performed over the blue half of the data, such that our measured pixel shifts would be independent of the observed fringes.

2.2. Telluric Contamination

Telluric contamination can be seen in a photometric G750M time series of the Na core, where we see a large variation of 4%–5% over an HST orbit (see Fig. 2). This variation is observed to be correlated with the HST orbital phase, repeats in a similar manner within each orbit of a visit, and occurs in all three visits analyzed. The contamination occurs for both Na lines over the same radial velocities, approximately −5 to −20 km s\(^{-1}\) from the Na D line cores, and is consistent with the expected velocity of the Earth during the HST observations. Given the size of the variation, it is too large to be due to the known instrumental systematic effects, which only affect the light curves ~0.1% during an orbit. In addition, other strong stellar lines such as Fe\(^{1}\) do not show any signs of similar contamination, ruling out possible detector efficiency and count-rate-dependent issues as the cause of the observed line core variations. The contamination can also be seen directly in the G750M spectra (Fig. 1), observed as an ~0.3 Å region in the Na D cores with little to no flux drop during planetary transit, although the actual contaminating region could result from a smaller unresolved region. Absorption by interstellar Na could also be affecting the contaminated region, although its effect would be constant and should not vary with the HST orbital period. The G750L data set also shows telluric Na D contamination, with H\(\alpha\) showing a similar effect as well. With lower resolution, the telluric signature is seen in photometric

![Image of G750M spectra of the Na D2 (top) and D1 (bottom) line cores. Plotted are all of the out-of-transit spectra (black plus signs), as well as the in-transit spectra (gray crosses). The uncertainty in counts for each point is dominated by photon noise (approximately ±400 photoelectrons) and is smaller than the plotted symbol size. The first out-of-transit spectrum in the series is plotted in histogram mode to show the size of each pixel, 0.554 A. Plotted together, each individual spectrum’s Doppler and pixel shifts (accurate to ~0.1 pixel) effectively add up to better sample the PSF of the stellar absorption line profile. The ~1.5% difference in flux during transit can be seen in the flux difference between the two spectra. However, the region showing contamination from 5889.5 to 5889.7 Å and from 5895.5 to 5895.8 Å shows little to no difference in flux during transit. This region shows variations with the HST orbital period, which we attribute to terrestrial-based sodium contamination.](image-url)
time series measurements of the line cores with, as expected, a correspondingly lower amplitude than the medium-resolution data.

2.3. Sodium Core Absorption Signature

We performed a double differential photometric measurement, similar to that of Charbonneau et al. (2002), selecting 23 wavelength bands from \( \sim 4 \) to 100 Å centered on the Na D lines. In the original analysis, Charbonneau et al. (2002) used three different wavelength bands. We chose to use the sum of two bands, each centered on one component of the Na doublet for the two wavelength bands smaller than the Na D doublet separation. For each band, we integrated the spectral flux, producing photometric light curves. We then normalized each visit to their average out-of-transit flux. The two narrowest bands showed terrestrial contamination modulating on the HST orbital period, which we removed by fitting each visit of the HST phased out-of-transit flux with a fourth-order polynomial, applying the fit to the in-transit orbits, as was done in a similar manner for the instrumental systematic effects. The photometric curves were then subtracted from a comparison band composed of the wavelength regions between \( 5818-5843 \) Å and \( 5943-5968 \) Å, picked to be identical to the “wide” comparison band in Charbonneau et al. (2002). The relative planetary Na absorption during transit was then measured as performed for the photometric analysis, including telluric corrections for each pixel. The resulting differential spectra show two statistically significant positive absorption peaks at the Na D lines with each having a large negative absorption core where the terrestrial contamination is largest (see Fig. 4). The telluric corrections on the Na lines have a modest effect on the absorption values outside the two largest contaminated pixels, as the telluric perturbation originates in a small subpixel region, with nearby pixels affected through the \( \sim 2 \) pixel instrument resolution. Excluding the two highly contaminated pixels effectively removes the majority of the contamination. In a 4 pixel bin, the average value of the two line cores is consistent with the 4.4 Å band photometry measurement, indicating that both approaches to correcting the contamination give consistent values. In our final limb-darkened, corrected transmission spectra (see § 4.3), we ultimately choose to apply the telluric corrections over the Na region while also leaving out the two highest

Fig. 2.—G750M photometry of the contaminated region. Plotted are photometric curves from the comparison region, which includes wavelengths between the Na lines \( 5818-5843 \) and \( 5943-5968 \) Å (black symbols), along with (top) 1 pixel between the Na line at 5893.4 Å that is completely uncontaminated. The plot also shows (middle) the same comparison region along with a highly contaminated pixel in the Na D1 line core at 5895.65 Å and (bottom) a nearby D1 line-core pixel that is largely uncontaminated. A representative error bar for the single-pixel photometric light curves is also shown in the lower left-hand corner; the error for the comparison region is significantly smaller than the symbol size. The three colors (purple, blue, and magenta) distinguish the three different HST visits, each consisting of five consecutive orbits (four orbits plotted). The telluric contaminated Na core can be seen as an \( \sim 4\%-5\% \) change in flux which modulates on the HST orbital cycle, repeating each orbit.
of D2/D1

et al. (2002) measurements (points) show that the Na D2 line is stronger than the absorption of the D1 line, with absorption values of 0.95% ± 0.02% and 0.051% ± 0.02%, respectively, and an absorption ratio of D2/D1 = 1.9 ± 0.9. The differential photometric measurement includes these two pixels, as their effects are negligible at large bandwidths and correctable for the 4.4 and 8.9 Å bands.

2.4. Limb-darkening Corrections

Correcting for limb-darkening effects allows a full transmission spectrum to be produced over the entire wavelength range of the STIS data. Similar to Ballester et al. (2007), we compute nonlinear limb-darkening coefficients to estimate and correct for the limb-darkening effects in the transit light curve. Although other limb-darkening laws could, in principle, be used (e.g., linear, quadratic, square root, or logarithmic), Claret (2000) found that the nonlinear law best represents the intensity distributions calculated from the one-dimensional plane-parallel atmospheric models used here.

We choose a Kurucz ATLAS stellar atmospheric model based on the model’s limb-darkening performance, its performance versus other one-dimensional models, model availability, and development within the literature. Phoenix models (Claret 2000) were found to predict a stronger limb darkening, compared to ATLAS models, by a few percent in the photometric bands B, V, R, I, J, and H but up to 40% weaker in the U and u’ bands. As the ATLAS models performed better compared to the Sun (see § 3.2), they were chosen over the Phoenix models for our study.

Most studies fitting transit light curves have chosen to either adopt an ATLAS stellar atmospheric model to account for limb-darkening effects or fit for limb-darkening parameters assuming a limb-darkening law. Knutson et al. (2007), who also used an ATLAS model for HD 209458, found similar planetary parameters when comparing the two methods, suggesting that model-based limb-darkening fits can perform well. Knutson et al. (2007) chose to use the limb darkening calculated from the ATLAS models in their results, due to the more precise planetary parameters found and concerns that fitted limb-darkening coefficients could be affected by residual correlation in the light curve. Such concerns are valid, as we found that residual correlated “red” noise ultimately limits the maximum signal-to-noise ratio (S/N) achievable in this study (see § 4.2). In addition, as our limb-darkening corrections can be on a pixel-by-pixel level, a light curve generated by any one particular pixel will not have sufficient S/N to accurately fit for limb-darkening coefficients.

The four nonlinear coefficients were computed using intensity spectra from a Kurucz one-dimensional plane-parallel model atmosphere calculated at 20 different angles spread uniformly in μ, fitting for the coefficients in the desired wavelength range. We calculated coefficients for the wavelength ranges used in the G750M differential measurements and each pixel of the G750M and G750L dispersions. Using the fit coefficients, we then used the theoretical transits of Mandel & Agol (2002) to determine and correct for limb-darkening effects. The best-fit system parameters of Knutson et al. (2007) were used to calculate the impact parameter and planetary/star radius contrast. For the differential photometric measurement probing the Na line cores, designed by Charbonneau et al. (2002) to be largely independent of limb-darkening, all the corrections had a less than 1 σ effect, with corrections to our narrowest band only increasing the absorption by 0.005% (see Fig. 3).

3. LIMB-DARKENING TESTS

3.1. Stellar Line versus Continuum Limb-darkening Strength Measurement

When measuring Na in the G750M photometric bandpass, limb-darkening effects increase with narrowing Na bands. Specifically, the stellar limb darkening is less in the Na line core compared to the surrounding continuum, producing transit light...
curves which are slightly more “box-shaped,” with a slightly smaller depth at midtransit. When comparing the transit light curves of narrow regions centered on the Na line to that of a surrounding continuum, the overall effect is to inherently underestimate the signature. Using the G750M spectra, we measured the magnitude of this effect using the stellar absorption lines (Fe i, Ca i, Si i, etc.) and the continuum. No other large planetary atmospheric signature other than Na was detected in the stellar lines of the G750M data (Charbonneau et al. 2002), making them excellent probes of the effects of limb darkening independent of differing planetary radii. We produced two separate photometric light curves, one that contained only the flux from the stellar continuum and one that contained only the flux from the stellar absorption lines. The Na region was not included in either photometry. We then normalized the two curves by their out-of-transit flux and took their difference. The average in-transit difference was measured, as in the differential Na analysis, to be $-0.0044\% \pm 0.0015\%$. Using the model atmosphere and calculating theoretical transits for the selected wavelengths, we found a theoretical difference of $-0.0038\%$, which matches the observed value. This confirms that (1) the limb darkening in the stellar lines is smaller than the continuum, (2) the effect is to reduce a planetary absorption signature when measuring within a stellar line and comparing to a nearby continuum, (3) the effect is small for a typical stellar absorption line (approximately $-0.004\%$), and (4) the strength of the effect is consistent with theoretical one-dimensional stellar atmosphere predictions. Although the one-dimensional stellar model used in our analysis does not contain a stellar chromosphere and assumes local thermodynamic equilibrium, the similarity between the model and data suggests that their effects are negligible on the transit light curve at these wavelengths and resolutions.

### 3.2. Performance of Solar-like ATLAS Limb-darkening Models

To gauge the performance and limitations of ATLAS atmospheric models in relation to analyzing planetary transits, we compared the limb-darkening predictions of a solar ATLAS model to the measured values of the Sun. The Sun represents the only test case where sufficiently high-precision optical limb-darkening data exist. HD 209458 is a solar-like star classified as a G0 V, while the Sun has a similar spectral type and is a G2 V. In analyzing the performance of model atmospheres, Bertone et al. (2004) found that the ATLAS models provide a good-fitting accuracy for the physical parameters, namely, $T_{\text{eff}}$ and $\log g$, for early-type main-sequence stars, B to F, down to late-G/early-K stars. For late-K and M stars, the models produce systematically poorer fits, due largely to the incomplete treatment of molecular opacity, which becomes important in cooler stars. As HD 209458 is slightly warmer than the Sun, molecular opacity should correspondingly play a slightly weaker role, and the performance of the ATLAS model itself should be similar for such main-sequence stars. Testing a solar ATLAS model versus solar data represents a test of the model itself and the assumptions within the model, as the physical parameters of the Sun, which are input into the atmospheric model, are both accurately and precisely measured. The uncertainty in the measured stellar parameters themselves, and its effect on limb-darkened, corrected transit light curves, is a separate issue dealt with in § 4.1.

Figure 5 shows high-precision solar limb-darkening measurements for 23 precisely chosen continuum wavelengths between 3000 and 11000 Å from Neckel & Labs (1994; hereafter NL94) and the corresponding solar ATLAS predictions, taken from the nonlinear solar ATLAS monochromatic coefficients computed by Claret (2000). The ATLAS model was calculated at a low resolution, so those wavelengths that matched closest were chosen. The ATLAS model performs well over $\mu$ values between $\sim 0.2$ and 1.0, slightly overestimating the strength of limb darkening consistently at every wavelength by a couple of percent (see Fig. 6). Our ATLAS solar comparison agrees with earlier studies (Castelli et al. 1997). The general overestimation by ATLAS models is expected, due to the inherent limitations of one-dimensional models when describing purely three-dimensional turbulent effects. This overestimation has been observed in both an F5 V and a K1 V star when comparing ATLAS models and three-dimensional hydrodynamical models to high-precision interferometric measurements (Aufdenberg et al. 2005; Bigot et al. 2006). The largest differences between the model and data exist at the very limb, where ATLAS models predict a dramatic increase.
in limb-darkening strength, a trend not observed in the solar limb-darkening data.

3.3. Simulated Solar Transits

The effect of adopting an ATLAS model to describe limb darkening in transit light curves can be seen by simulating transits with the high-precision NL94 solar data and comparing them to transits using an ATLAS solar model. For this purpose, we modeled transit light curves using the best-fit parameters of the HD 209458 system and both of the nonlinear limb-darkening coefficients predicted by the ATLAS solar model, as well as the solar data (see Fig. 7). With overall stronger limb darkening, the ATLAS models produced systematically deeper transits between second and third contact and shallower light curves during ingress and egress. The largest differences between the solar and ATLAS transits occur during ingress and egress, where transit parameter fits are relatively insensitive to the planet-to-star radius contrast, $R_{pl}/R_{\text{star}}$ which is the desired parameter for this study. This insensitivity, combined with its small integrated flux contribution of the limb in contrast to the rest of the stellar surface, limits the overall effect of the observed large limb model deficiencies during midtransit. At midtransit, an $\sim 2\%$ overestimation in limb darkening (corresponding to $\mu \sim 0.5$) translates to an extra transit flux drop of $\sim 0.01\%$. These trends are seen to be largely reproducible, wavelength-to-wavelength, at a $\pm 0.005\% - 0.006\%$ level. The near-UV (NUV), where the strength of limb darkening is larger, also exhibits this general overestimation trend, but with a slightly larger scatter compared to the rest of the optical.

We corrected the transit light curve, computed with the NL94 data, for limb-darkening effects as predicted by the ATLAS model. This procedure is comparable to our limb-darkening correction procedure adopted for HD 209458b, i.e., using the predicted limb darkening from ATLAS models to correct transit light curves which are the result of the actual limb darkening of a star. These corrected light curves can be seen in Figure 8 (top curves), where the limb-darkening overestimation by the ATLAS models translates into a slightly lower flux ratio at midtransit and slightly higher flux ratio just before third contact or just after second contact, with this general trend seen at all wavelengths. Thus, introducing the ATLAS model produces two systematic effects: (1) a wavelength-to-wavelength flux scatter of up to $\pm 0.005\% - 0.006\%$ and (2) a phase-dependent trend of overcorrecting the flux during the middle approximately two-thirds of in-transit times and undercorrecting during the remaining one-third. The first systematic effect is likely due to the internal accuracy of the ATLAS models, where uncertainties in opacity database tables and metal line blanketing, for instance, could reveal themselves. The second systematic effect is clearly related to the general trend of ATLAS models overpredicting the strength of limb darkening at a similar level in all wavelengths and tied to the treatment of convection in one-dimensional atmospheric models.

To test whether the broadband NUV and Na absorption features seen in our data could successfully be recovered using ATLAS model limb-darkening corrections, we also simulated transits using the NL94 data, which contained an extra $0.01 R_{\text{pl}}$ in planetary radii for the NUV and an extra $0.02 R_{\text{pl}}$ in planetary radii for selected Na-like wavelengths. These simulated absorption transit light curves were then corrected with those generated with solar ATLAS models and the best-fit planetary radius (see Fig. 8). The two simulated absorptions can easily be identified above that of the zero-absorption case, with the general phase-dependent trend largely preserved. Therefore, when averaging over a series of in-transit spectra and comparing absorption values at different wavelengths, the values of the relative absorption differences will largely be retained after the ATLAS corrections, but the flux level itself could change slightly, depending on the phase sampling (effect 2). At the extremes, using only data at midtransit, second contact, or third contact, this absolute flux difference is seen to be up to $\sim 0.01\%$. These trends are observed in our absorption simulations. When we average over the in-transit phases and recover the input absorption signal (see Fig. 9), we get relative absorption values (above the ATLAS-corrected zero-absorption case) of $0.026\% \pm 0.003\%$ and $0.046\% \pm 0.002\%$ for the NUV-like and Na-like signatures, respectively. These values closely match the input $0.022\%$ and $0.044\%$ values, indicating that ATLAS model corrections can retain relative absorption signals at a level of around $\pm 0.003\%$. Selecting subsets which contained only similar orbital phases produced equivalent...
differential results, indicating that effect 2 has a limited impact on relative absorption measurements.

3.4. Application of Solar Test to HD 209458

The residual G430L and G750L ATLAS-fit light curves (in addition, see Fig. 4 of Knutson et al. 2007) show the same trends as observed in our solar test. Namely, the HD 209458 transit light curves show a systematically slightly higher residual flux at midtransit compared to second and third contact at an \( \approx 0.01\%-0.02\% \) flux level, as in our solar test, with this trend repeated at all optical wavelengths. This indicates that the performance of the ATLAS models is similar for both HD 209458 and the Sun. When fitting for the planetary radius, in the case of HD 209458, adopting an ATLAS limb-darkening model would therefore seem to systematically fit a slightly smaller \( R_p/R_\text{star} \) ratio at every wavelength and thus result in smaller planetary radii (given uniform and complete phase coverage). However, this systematic error is much smaller than the uncertainty due to the stellar mass and has a negligible effect on the final determined planetary radius. This is reflected in the analysis of Knutson et al. (2007), who find planetary radii values that differ by less than 1 \( \sigma \) when using either the method of adopting ATLAS models or fitting for limb-darkening coefficients.

Our analysis shows that, over the entire optical regime 3000–10000 Å, it is possible to use ATLAS models to accurately measure the relative value of planetary radii to around a 0.005%-0.006% precision level in solar-type stars. This precision level is comparable to our highest S/N achieved when averaging over large wavelength ranges (>1000 Å) and multiple exposures (see § 4.2). Thus, the systematic errors introduced by assuming ATLAS models are negligible for our purposes of probing relative transit depths for signals >0.01%. The success of using ATLAS models in analyzing transits of the K0 V star HD 189733 (Pont et al. 2008) further highlights their capabilities. As a relatively cooler K star, molecular opacity plays a correspondingly larger role in HD 189733 than the Sun or HD 209458, yet an ATLAS model provides superior fits compared to leaving the limb-darkening coefficients as free parameters in transit fits (F. Pont 2008, private communication). As more high-precision measurements become available across a wider wavelength range, care will have to be taken when comparing transit parameters determined by separate analyses, which make different limb-darkening assumptions.

4. ANALYSIS

4.1. Stellar Parameter Uncertainties

Uncertainty in the stellar parameters \( (T_\text{eff}, \log g, \log_{10}[M/H], v_\text{turb}) \) can lead to a range of possible stellar models and thus a range of predicted limb darkening. As limb darkening derives mainly from intensity differences between blackbodies at different temperatures, changes in the \( T_\text{eff} \) of the star have the largest effect on the calculated limb darkening. For HD 209458, the range of \( T_\text{eff} \) found in the literature is from \(~6000\) to \( 6100 \) K (6030 \( \pm \) 140 K, Allende Prieto & Lambert 1999; 6000 \( \pm \) 50 K, Mazeh et al. 2000; 6117 \( \pm \) 26 K, Santos et al. 2004; 6099 K, Fischer & Valenti 2005), and we have adopted a \( T_\text{eff} \) of 6100 for this study.

We estimate the uncertainty of limb darkening inferred from stellar parameter uncertainty by estimating how large a temperature increase or decrease is necessary to produce a difference flux of 0.02% at midtransit between bandpasses, which is near the absorption signal detected in the NUV region. In our wavelength regime, the flux at midtransit in the NUV is the most sensitive to changes in limb darkening, making this test an upper limit to the \( T_\text{eff} \) change needed at longer wavelength bandpasses and other phases. From Claret (2000), we use nonlinear limb darkening for the photometric bandpasses \( U \) and \( B \) calculated at three different temperatures and measure the transit depth that results. The \( U \) and \( B \) wavelength ranges roughly correspond to the \( \sim 0.03\% \) flux difference observed in our transmission spectrum between the NUV (3000–3700 Å) and nearby optical (4000–5500 Å). Using the best-fit HD 209458 parameters, we calculate that a \( T_\text{eff} \) change from 6000 to 5500 K changes the \( U \) band midtransit depth by 0.05% and the \( B \) band by 0.03%, resulting in an \( \sim 0.02\% \) change in relative flux depth. Likewise, changing the \( T_\text{eff} \) from 6000 to 6500 K changes the midtransit depth by 0.05% in the \( U \) band and 0.03% in the \( B \) band, also resulting in an \( \sim 0.02\% \) change. Thus, an uncertainty of only \( \pm 50 \) K for the \( T_\text{eff} \) of HD 209458 corresponds to a NUV limb-darkening uncertainty of only \( \sim 0.002\% \), a level well below our S/N level.

4.2. Red Noise

We quantify the level of systematic “red noise” in our transmission spectrum following the procedures detailed in Pont et al. (2006). In the absence of red noise, averaging over successive measurements increases the S/N in a predictable manner, with the standard deviation of the average varying inversely with the square root of the number of measurements, \( N_t \), and the variance and standard deviation from a distribution of points binned by \( N_t \) and calculate the variance and standard deviation from a distribution of points binned by \( N_t \) and successive measurements. This is repeated for an increasingly larger number of \( N_t \) and \( N_w \) bin sizes. We check different binning combinations to ensure that correlated noise does not appear at our binning frequency. Plotted in Figure 10 is a sample of red noise estimation from the G750L data in the 6200 Å region. Typically, photon noise is seen to dominate down to precision levels of \( \sim 0.02\% \), where the effects of red noise start to become visible. For most of the optical, binning...
over large wavelength regions and multiple exposures allows us to reach precision levels of ~0.006%, although after binning beyond ~30 measurements, the S/N usually does not significantly increase.

As binning over larger wavelength bands has diminishing returns, the presence of red noise gives our transmission spectra a S/N level at moderate wavelength bins similar to that at larger wavelength bins. This is reflected in Figure 11, where all five 1 σ error bars calculated to show the significance of observed broadband features are similar to each other, although they differ in wavelength breadth. The broadband errors are also similar to the typical 1 σ uncertainty levels when binning over 16 pixels, 0.011%.

4.3. Broadband Absorption Signatures and Optical Transit Transmission Spectra

Using the G750M and G750L spectra, we used a limb-darkened corrected transit spectral ratio to probe the full wavelength dependence of atmospheric absorption features, producing an optical transmission spectra. The relative planetary radii, when comparing between different wavelengths, are largely independent of the uncertainties of the stellar parameters and can be determined much more precisely than the actual radius values themselves (Knutson et al. 2007). Putting the low-resolution G750L data together with the G430L data (Ballester et al. 2007) allows a probe of the region from 3000 to 10000 Å, with the Na D lines located in a high-S/N region in the middle.

Our low-resolution transmission spectra were produced by interpolating all the systematic error-corrected spectra onto a common wavelength scale. We then produced both an average out-of-transit spectrum and an average limb-darkened corrected in-transit spectrum (using phases between second and third contact), with the ratio of the two providing our ratio spectrum. The error was computed by taking into account both photon noise and red noise, as outlined in § 4.2. The G750L spectral region blueward of 5548 Å at the edge of the spectrum was omitted, as it contains a region of very rapidly decreasing response whose spectral ratio value is 0.03% lower when compared to the higher S/N overlapping the G430L region. This omitted G750L region is seen to have a larger amount of red noise, limiting the accuracy levels in the region to ~0.02%–0.03%, which likely explains the flux difference between the two data sets. We used the best-fit value of $1 - (R_{pl}/R_{star})$ ~ 1.455% from Knutson et al. (2007) as our baseline transit depth reference, subtracting our spectral ratio from that value to construct a plot of the atmospheric absorption profile (Fig. 11). We also performed the same procedure on the G750M data, although a shift of +0.00023 was needed to generate an average absorption level consistent with that of the low-resolution data. As the medium-resolution data were taken 3 yr before the low-resolution data, long-term differences in starspots, for instance, could easily account for the small measured difference in absolute absorption level. The effects of the uncertainties in the transit parameters were tested by varying the period, inclination, stellar radius, and stellar mass by 1 σ (as reported in Knutson et al. 2007) and producing a new transmission spectrum. Due to limb darkening, changing the inclination has a larger effect than the other parameters, although all produce negligible effects on the final transmission spectrum. For instance, even for the NUV at 3000 Å, which contains strong limb darkening, changing the inclination by 1 σ still only has a 0.002% effect on the final transmission spectrum.

The wavelength range of the G750L grating extends in the red past 1 μm, although with significant detector fringing (Knutson et al. 2007). The data reduction procedures used here were insufficient to correct for the large fringing effect on a pixel-by-pixel basis. Consequently, those regions which showed signs of significant fringing were left out of this analysis, corresponding to wavelength regions longer than ~8000 Å.

5. DISCUSSION AND CONCLUSIONS

Our optical transit transmission spectrum combining all three gratings (see Fig. 11) gives a consistent and more complete view of the observed absorption features present in HD 209458b. The low-resolution transmission spectrum reveals three prominent
broadband absorption features: (1) a NUV absorption below \(~4000\) Å, (2) a strong broadband Na feature, and (3) an additional absorption feature at \(~6250\) Å. The medium-resolution data (1) show the strong Na core absorption (\(~0.065\)% in a 4.4 Å line-core band) and (2) confirm the presence of the absorption feature at \(~6250\) Å. The NUV absorption region was first reported by Ballester et al. (2007) and is interpreted, along with wavelengths shorter than 5000 Å by Lecavelier des Etangs et al. (2008), to be Rayleigh scattering by H2, while the long-wavelength continuum redward of \(~6200\) Å can be explained by absorption by TiO and VO (Désert et al. 2008).

Seen at low resolution, the wide Na wings and additional absorption at \(~6250\) Å place the medium-resolution data on an apparent plateau, \(~0.045\)% above that of the minimum absorption levels seen around 5000 Å and redward of 7000 Å. As such, the total Na core absorption reaches to \(~0.11\)% within a 4.4 Å band above the minimum level at 5000 Å. This strong Na absorption is much closer to the original cloudless predictions (Seager & Sasselov 2000; Brown 2001; Hubbard et al. 2001), albeit with a more complicated Na line absorption profile (Sing et al. 2008). The large Na line-core absorption would seem to disfavor high-altitude clouds or haze cutting into the absorption signature, as recently seen in HD 189733b by Pont et al. (2008). The distinct transition from broad Na features at low absorption depths to sharp, narrow features at high absorption depths, however, indicates that condensation of Na is likely in the upper atmosphere, which would deplete those higher altitudes of atomic Na (Sing et al. 2008).

The complicated Na D line shape and telluric contamination illustrate some of the difficulties in observing this atmospheric absorption signature in HD 209458b. Ground-based spectroscopic measurements, such as the Na detected in HD 189733b (Redfield et al. 2008), require air-mass corrections and are only sensitive to the Na D line cores where the atmospheric atomic Na in HD 209458b is substantially depleted. Furthermore, telluric Na contamination has a nonnegligible impact on the signature, further masking the small signal.

As seen here with Na, transit spectra at multiple resolutions are needed to properly interpret complex transmission features. Data at low resolution, covering a broad wavelength range, place the overall features in context, while data at medium or high resolution are better suited to identify specific absorption features and probe planetary atmospheres over a large pressure range. Our tests show that for exoplanets with solar-type stellar hosts, such transmission spectra can be corrected for limb-darkening effects to a high precision with current atmospheric models.

Space-based observations on exoplanetary transits offer the extraordinary ability to provide exoplanetary transmission spectra at multiple resolutions, leading to detailed atmospheric composition information. As the S/N, resolution, and wavelength coverage of such spectra improve, multiple species can be detected and atmospheric properties inferred. The data quality of the spectra used here and results herein help indicate the full potential of the transit method, which can be further exploited with the impending repair of HST and future launch of the James Webb Space Telescope.

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