Signatures of relativistic protons in CTA blazar spectra

ZECH, A.¹, CERRUTI, M.¹² FOR THE CTA CONSORTIUM

¹ LUTH, Observatoire de Paris, CNRS, Université Paris Diderot ; 5 Place Jules Janssen, 92190 Meudon, France
² now at the Smithsonian Astrophysical Observatory, Cambridge, MA, USA

andreas.zech@obspm.fr

Abstract: The Cherenkov Telescope Array (CTA) will provide very-high-energy (VHE; ≳30 GeV) γ-ray spectra of unprecedented resolution over a large energy range. It is likely that, at least for bright sources, spectral features will be revealed that cannot be distinguished with the current generation of Cherenkov telescopes. We are investigating the capability of the CTA to detect spectral signatures in TeV bright BL Lac objects that could reveal the emission from a population of relativistic protons in those sources. Such a detection would be crucial for the long-standing question of the origin of extragalactic cosmic rays. As a first step, the expected VHE emission from hadronic scenarios is compared to the spectral features that might arise in the more commonly assumed leptonic scenarios for a given source. We also evaluate the impact of different array configurations on the detectability of such features.

Keywords: VHE gamma rays, AGN, emission models

1 Introduction

The double-bumped spectral energy distributions (SEDs) of blazars are in general very well represented with simple emission models that ascribe the lower energy bump (in the optical to X-ray range) to synchrotron emission from electrons or electron-positron pairs and the higher energy bump (peaking in γ-rays) either to synchrotron self-Compton (SSC) emission of the same population of leptons or to external inverse Compton (EIC) emission with external photon fields. While the former scenario is very successful for the description of high-frequency peaked BL Lacs (HBLs), the addition of external fields, from the accretion disk, broad-line region or dust torus, is indicated in flat-spectrum radio quasars (FSRQs) and probably in low/intermediate-frequency peaked BL Lacs (LBLs/IBLs). These leptonic models have the advantage of providing a simple explanation, often based on a single homogeneous emission region, for the observed SEDs over a very large energy range. They also reproduce the characteristic rapid variability of blazars, except in the most extreme cases [1] [2]. An example for an SSC interpretation of the SED from an HBL is shown in Fig. 1.

A more complex alternative to leptonic scenarios is given by hadronic emission models that also ascribe the bump at lower energies to electron synchrotron emission, but suppose hadronic interactions with magnetic or photon fields to be the main cause for the high-energy emission. Such scenarios imply a direct link between the electromagnetic spectrum from a source and an accompanying flux of ultra-relativistic cosmic rays and neutrinos. Their study is thus of high interest for the question of the origin of ultrahigh energy cosmic rays (UHECRs; energy ≳10¹⁸ eV), although the currently available blazar data do not allow us to distinguish them from their more “economic” (in terms of the number of free parameters) leptonic counterparts. In hadronic scenarios, it is generally more difficult to ac-

Figure 1: SED of PKS 2155-304 in 2008 [3] with SSC model.

Figure 2: SED of PKS 2155-304 in 2008 [3] with two different hadronic model curves. The solution “hadro 1” shown on the left panel uses a magnetic field of 80 G; the solution “hadro 2” on the right panel uses a lower magnetic field of 50 G and a higher proton density.
count for very rapid variability due to the larger masses of hadrons, which imply longer time scales for acceleration and radiative cooling. This however depends on the physical conditions inside the source and certain scenarios achieve very short time scales (e.g. [3]). An additional problem is that energy requirements are in general larger for hadronic than for leptonic scenarios [5]. Fig. 2 shows two different hadronic interpretations of the same low state of the HBL PKS 2155-304 as in Fig. 1.

The current generation of Cherenkov telescopes has played a key role in the investigation of Galactic cosmic-ray sources. The great improvements in energy coverage, sensitivity and energy resolution expected from the CTA [6, 7], currently in its prototyping phase, may provide us with the tool needed to break the ambiguity between leptonic and hadronic scenarios for extragalactic sources, leading to new insights in the origin of UHECRs. We want to discuss here which spectral signatures could betray the presence of relativistic protons among the γ-ray emitting particles in blazars.

2 Brief description of our lepto-hadronic code

We have developed a stationary lepto-hadronic code that allows the user to switch between a standard SSC scenario and a hadronic scenario by changing the parameters that describe the conditions in the source region and the leptonic and hadronic particle populations. This global approach gives also access to mixed — lepto-hadronic — scenarios, where emission from both primary leptons and hadrons contribute to the high-energy bump [8].

The leptonic part of the code is based on a stationary one-zone SSC model [9]. This “blob-in-jet” model employs an electron distribution parametrized by a broken power law, interacting with a homogeneous, tangled magnetic field. A full description of internal photon-photon absorption, the γ-ray emission from secondary pairs and the SSC emission in the Thomson and Klein-Nishina regimes is given. This code was successfully applied to SEDs of several HBLs (e.g. [10], [11], [12],...).

The description of the hadronic interactions is based on the assumption that electrons and protons are co-accelerated in the same emission region to reduce the number of free parameters. The maximum proton energy is limited by a comparison of the adiabatic and radiative cooling time scales with the acceleration time scale. In addition, the gyroradius of the most energetic protons is constrained by the size of the emission region. The code calculates proton synchrotron emission and proton-photon interactions on the synchrotron photons of the primary electron population. Proton-meson production is modeled using the detailed Monte Carlo simulations provided by the SOPHIA code [13], while Bethe-Heitler pair production is based on an analytic formulation [14]. The use of SOPHIA permits the detailed simulation of muon synchrotron emission, which can be a very significant component in the description of BL Lac spectra. Decaying pions and other mesons, generated in large numbers in proton-photon interactions, inject very energetic leptons and γ-rays into the emission region, which trigger synchrotron-pair cascades when interacting with the low-energy photon field. We follow these cascades for several generations to arrive at a steady state. When calculating spectra of muons and cascade particles, the particle distributions are cooled following the treatment in [15] considering radiative losses by synchrotron and Inverse Compton cooling, as well as adiabatic losses. Proton-proton interactions are considered negligible in the scenarios under study.

The VHE photon flux, generated in the leptonic or hadronic scenario, is absorbed on the internal synchrotron photon field and on the extragalactic background light (EBL) between the source and the observer. The model by [16] is used by default to describe the EBL, but other models can be chosen as well.

3 Spectral signatures from hadronic emission models

Even if one was to admit that the detection of (very) rapid variability of the high energy flux during flares, from sources such as PKS 2155-304, Mrk 421 or Mrk 501, etc., favours lepton blazar emission scenarios, this does not exclude a significant contribution from hadrons, especially during low states or for HBLs with very hard spectra, which show little sign of variability [17]. It is thus promising to search for signatures in the emission spectra of blazars that might be specific to hadronic scenarios.

Such a signature can be seen in Fig. 2, the occurrence of synchrotron-pair cascades in the emission region adds a hard component to the very high energy spectrum that should be visible as spectral hardening in the TeV range. The appearance of such a “cascade bump” in the TeV spectrum seems to be a general feature of hadronic models (e.g. [18], [19]). The luminosity of the cascade and muon synchrotron emission compared to the proton synchrotron emission depends on the ratio of the primary particle density in the source compared to the magnetic field strength. Increasing the proton density while lowering the magnetic field leads to a stronger contribution from the products of proton-photon interaction and an earlier onset of the “cascade bump”. This can be seen in Fig. 2 where the right panel shows a hadronic scenario with a lower magnetic field and more prominent contribution from secondary particles than the scenario in the left panel.

4 A case study

To start investigating the detectability of the “cascade bump” signature in blazar spectra, we will focus here on a single test case and try to distinguish the leptonic and hadronic scenarios for the SED of PKS 2155-304 in 2008 shown in Fig. 1 and 2.

We have simulated spectra as they would be seen with CTA for different array configurations and exposure times, assuming first the leptonic and then the hadron scenarios. The model curves are taken as intrinsic source spectra and spectral points are calculated based on simulated performance curves (effective area, energy resolution, angular resolution, background rate) for a given CTA configuration [25]. As an example, Fig. 3 shows simulated spectra for the leptonic and the “hadro 1” scenario.

A simple logparabolic fit of the form \(dN/dE = N_0(E/E_0)^{-\alpha} \exp(-\delta \log(E/E_0))\) was found as a good means to discriminate the leptonic and hadronic models, in this case. The simulated spectra were fitted above 100 GeV, resulting in an optimum fit probability for the given SED. To account for statistical fluctuations in the CTA spectra due to photon
As expected, the discrimination between the leptonic and hadronic scenarios is no longer distinguishable. Observations of a single source in its low state over a relatively long time period would thus be required for the distinction of hadronic spectral features from a stationary source.

4.1 Influence of the EBL model

To evaluate the influence of the EBL model on the above results, we have replaced the description by [20] with the model by [21], yielding a lower level. As expected, the discrimination between the leptonic and hadronic curves suffers from a stronger absorption on the EBL, which has a tendency to wash out spectral features. Better discrimination is achieved if one assumes a low EBL level. While the different models lead to more or less overlap between the probability distributions, the threshold between the hadronic and leptonic distributions remains at about the same probability.

4.2 Influence of the exposure time

For longer exposure times, the “cascade bump” is more clearly visible due to the reduced impact of statistical fluctuations, which helps to distinguish spectral features. When reducing the exposure time to 5 h (Fig. 5), the probability distributions for the leptonic and hadronic scenarios are no longer distinguishable. Observations of a single source in its low state over a relatively long time period would thus be required for the distinction of hadronic spectral features from a stationary source.

4.3 Influence of the CTA array configuration

We have also repeated the study of our test case for different array configurations, which were proposed in the CTA Design Study to optimize the mixture of different telescope types and their arrangement within the array. Different configurations result in different sensitivity curves, angular and energy resolution. Table 1 shows the amount of overlap between the leptonic and hadronic probability distributions for different configurations for the two hadronic scenarios.
solutions shown in Fig. 2. The overlap (in % out of 200 trials) can be used as a rough indicator for the (lack of) discriminative power of the logparabolic fit.

| config. | telescopes | hadro 1 | hadro 2 |
|---------|------------|---------|---------|
| B       | 5 LST, 37 MST | 0       | 9       |
| C       | 29 MST, 26 SST | 0       | 4       |
| E       | 4 LST, 23 MST, 32 SST | 5       | 19      |
| I       | 3 LST, 18 MST, 56 SST | 3       | 13      |
| NA      | 4 LST, 17 MST | 10      | 35      |
| NB      | 3 LST, 17 MST | 12      | 27      |

Table 1: Discrimination between the SSC and hadronic models for different CTA array configurations. First two columns: array configuration and telescope content (LST: Large-Size Telescopes, MST: Medium-Size Telescopes, SST: Small-Size Telescopes); third and fourth column: overlap (in %) for 50 h exposure time between the SSC model and the hadronic models.

For both hadronic solutions in our test case, the Northern Array configurations “NA” and “NB”, seem to perform worst. The “cascade bump” in the “hadro 2” solution is statistically more difficult to distinguish from the SSC curve for all configurations, compared to the “hadro 1” solution. It can be seen that the difference in the probability distributions for two different hadronic solutions for the same SED is already quite important. It will thus be difficult to draw a more general conclusion from such studies regarding the preference for a certain array configuration. More sources need to be studied in a systematic way for this purpose. A preliminary evaluation of a hadronic interpretation of the low-state SED of Mrk 421 tends to favour configurations with a well balanced energy coverage (“E”, “I”).

5 General Considerations

In the one-zone framework, the “cascade bump” is a unique signature for hadronic models, if external photon fields and second-order SSC emission can be neglected. For HBLs, this is generally assumed to be the case. For other types of blazars, where photon fields, e.g. from the broad line region, seem to play a more important role, spectral “bumps” in the TeV range might arise from photon-photon absorption [22] or from additional external Compton components.

It can also not be excluded that a hardening in the high energy spectrum is caused by a second emission region, populated with an electron distribution with a harder spectrum than the dominant region. Emission from such a second zone might be hidden below the synchrotron and SSC bumps of the first zone and emerge only at TeV energies. However, for very hard blazar spectra, a potential contribution from such a second emission region might be sufficiently suppressed due to Klein-Nishina effects to leave only the hadronic interpretation for spectral hardening.

There is also the possibility that external cascades, triggered by UHE protons or γ-rays that leave the source and interact with the EBL and CMB photon fields, contribute to the observed VHE emission. In these scenarios, spectral hardening can also arise in the TeV range and will need to be distinguished from the internal “cascade bumps” [23].

When trying to connect the VHE signatures to the question of the origin of UHECRs, given our still limited knowledge of the nature of these particles, one will also have to consider the effect of relativistic nuclei, rather than protons, inside the source. Since photo-disintegration is competing with photo-hadron interactions in the case of nuclei, one would expect a smaller contribution from synchrotron-pair cascades inside the source [24], but this should be studied in more detail.

6 Conclusions & Outlook

For the specific example studied here (PKS 2155-304 in a low state), the “cascade bumps” from two hadronic solutions have been found to be statistically distinguishable from a leptonic SSC scenario with a simple logparabolic fit to the VHE spectrum. A sufficiently long exposure time (~50h) is needed and the discriminative power depends also on the array configuration and on the assumed EBL density.

In this first study, we have used a relatively simple approach to distinguish hadronic from leptonic model curves. A more elaborate fit function and a fitting procedure using forward folding might arrive at a clearer discrimination for shorter exposure times. The performance curves used here to describe the different CTA configurations were determined with a simple “Hillas”-type reconstruction of simulated air showers, while more efficient methods are known to improve the sensitivity and energy resolution due to better event reconstruction and background rejection [25].

This would also translate into shorter exposure times for the discrimination between two given scenarios.

Acknowledgment: This work made use of the macros written by D. Mazin for CTA. We wish to thank C. Boisson, H. Sol, S. Inoue and D. Mazin for very helpful discussions. We acknowledge the support from the LEA ELGA, Observatoire de Paris, CTA and CNRS, and from the agencies and organizations listed in this page: http://www.cta-observatory.org/?q=node/22

References

[1] F. Aharonian et al. (HESS Coll.), ApJ Lett. 664 (2007) 71
[2] J. Albert et al. (MAGIC Coll.), ApJ 669 (2007) 862
[3] F. Aharonian et al. (HESS Coll.), ApJ Lett. 696 (2009) 150
[4] M. V. Barkov, F. A. Aharonian et al., ApJ 749 (2012) 119
[5] M. Sikora, proc. of the IAU Symposium 275, 2010
[6] B. S. Acharya et al. (CTA Cons.), 2013, A.Ph. 43, 3
[7] M. Actis et al. (CTA Cons.), Exp Astron 32 (2011) 193
[8] M. Cerruti, A. Zech et al., proc. of Gamma 2012
[9] K. Katarzynski, H. Sol et al., A&A 367 (2001) 809
[10] A. Abramowski et al. (HESS Coll.), A&A 539 (2012) 149
[11] A. Abramowski et al. (HESS Coll.), A&A 542 (2012) 94
[12] A. Abramowski et al. (HESS Coll.), A&A 554 (2013) astro-ph/1303.1628
[13] A. Mücke, R. Engel et al., Computer Physics Communications 124 (2000) 290
[14] S. R. Kelner & F. A. Aharonian, Phys. Rev. D 78 (2008) 034013
[15] S. Inoue & F. Takahara, ApJ 463 (1996) 555
[16] A. Franceschini, G. Rodighiero et al., A&A, 487 (2008) 837
[17] G. D. Senturk, M. Errando et al., ApJ 764 (2013) 119
[18] A. A. Abdo et al. (Fermi-LAT Coll.), ApJ 736 (2011) 131
[19] M. Boettcher, A. Reimer et al., 2013, astro-ph/1304.0605
[20] T. M. Kneiske, T. Bretz et al., A&A 413 (2004) 807
[21] T. M. Kneiske & H. Dole, 2010, A&A, 515, 19
[22] F. A. Aharonian et al., MNRAS 387 (2008) 1206
[23] H. Takami, K. Murase et al., astro-ph/1305.2138
[24] K. Murase, C. D. Dermer et al., ApJ 749 (2012) 63
[25] K. Bernlohr, A. Barnacka et al. A.Ph. 43 (2013) 171