Near-infrared observations of Type Ia supernovae: the best known standard candle for cosmology

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ABSTRACT

We present an analysis of the Hubble diagram for 12 normal Type Ia supernovae (SNe Ia) observed in the near-infrared (NIR) J and H bands. We select SNe exclusively from the redshift range 0.03 < z < 0.09 to reduce uncertainties coming from peculiar velocities while remaining in a cosmologically well-understood region. All of the SNe in our sample exhibit no spectral or B-band light-curve peculiarities and lie in the B-band stretch range of 0.8–1.15. Our results suggest that SNe Ia observed in the NIR are the best known standard candles. We fit previously determined NIR light-curve templates to new high-precision data to derive peak magnitudes and to determine the scatter about the Hubble line. Photometry of the 12 SNe is presented in the natural system. Using a standard cosmology of (H₀, Ω_m, Ω_Λ) = (70, 0.27, 0.73), we find a median J-band absolute magnitude of M_J = −18.39 with a scatter of σ_J = 0.116 and a median H-band absolute magnitude of M_H = −18.36 with a scatter of σ_H = 0.085. The scatter in the H band is the smallest yet measured. We search for correlations between residuals in the J- and H-band Hubble diagrams and SN properties, such as SN colour, B-band stretch and the projected distance from the centre of the host galaxy. The only significant correlation is between the J-band Hubble residual and the J − H pseudo-colour. We also examine how the scatter changes when fewer points in the NIR are used to constrain the light curve. With a single point in the H band taken anywhere from 10 d before to 15 d after B-band maximum light and a prior on the date of H-band maximum set from the date of B-band maximum, we find that we can measure distances to an accuracy of 6 per cent. The precision of SNe Ia in the NIR provides new opportunities for precision measurements of both the expansion history of the universe and peculiar velocities of nearby galaxies.

Key words: cosmology: observations – distance scale.

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1 INTRODUCTION

Type Ia supernovae (SNe Ia) are widely accepted to be excellent standardizable candles at optical wavelengths. The development of various empirical corrections, which can relate peak luminosity and light-curve shape (Phillips 1993; Hamuy et al. 1996), SN colour (Tripp 1998; Jha, Riess & Kirshner 2007), spectral information (Brongersma et al. 2002; Walker et al. 2011; Silverman et al. 2012) and host-galaxy mass (Sullivan et al. 2010), allows SNe Ia peak magnitudes to serve as distance indicators with a corrected scatter1 of as low as 0.13 mag (Conley et al. 2011; Silverman et al. 2012).

There has been growing evidence, though, that SNe Ia may be more accurate in the near-infrared (NIR) (for a recent review, see Phillips 2011). NIR light experiences less attenuation from dust than light at optical wavelengths, and theoretical models predict a smaller intrinsic dispersion in the NIR peak magnitude (Kasen 2006). Furthermore, there appears to be little or no relationship between light-curve shape and peak luminosity (Wood-Vasey et al. 2008; Folt et al. 2010; Mandel, Narayan & Kirshner 2011), meaning no empirical corrections need to be applied.

Recent studies (Meikle 2000; Krisciunas, Phillips & Suntzef 2004; Wood-Vasey et al. 2008; Folt et al. 2010; Mandel et al. 2011; Katter et al. 2012) have shown that SNe Ia in the NIR can be as reliable as corrected SNe Ia observed at optical wavelengths, but these studies have had some limitations. Foremost among these limitations is that the SNe were observed at distances that are not sufficiently large to place them in the Hubble flow (defined here as \( z > 0.03 \)) and hence are affected by peculiar velocities. Here we present a study of a sample of SNe Ia selected to lie exclusively in the redshift range \( 0.03 < z < 0.09 \) in order to minimize redshift uncertainty due to peculiar velocity. We use NIR images on 8-m class telescopes to observe SNe in our sample, resulting in photometry that is more precise than that achieved for the handful of SNe that have been observed beyond \( z = 0.03 \).

2 SNE IA SAMPLE

The SNe Ia used in our sample are summarized in Table 2. These SNe Ia were discovered by the Palomar Transient Factory (PTF; Law et al. 2009; Rau et al. 2009), spectroscopically confirmed, observed in the \( g, r, i \) bands (Maguire et al., in preparation) at the Liverpool Telescope and followed up in both the NIR \( J \) and \( H \) bands, each with four epochs of observation. PTF09dsc was observed using High Acuity Wide field K-band Imager (HAWK-I) (Casali et al. 2006) on the European Southern Observatory (ESO) 8.1-m Very Large Telescope (VLT), and all others were observed using NIRI,2 the NIR Imager and Spectrometer (Hodapp et al. 2000) on Gemini Observatory’s 8.2-m Gemini-North telescope. Reference images were taken \( \sim 1 \) year after initial observations so that host-galaxy light could be removed.

The data were processed in a standard manner using IRAF3 and our own scripts. Darks were used to remove the pedestal, and flats were used to remove pixel-to-pixel sensitivity variations. We used the XDMSUMP package in IRAF to remove the sky and a modified version of IPIS2 (Alard 2000) to subtract the reference image of each SN from the images with SN light, leaving just light of the SN. Zero-points were derived from standards found in the Persson et al. (1998) catalogue. All magnitudes are reported in the natural system of each instrument. \( K \)-corrections were calculated using the revised spectral template of Hsiao et al. (2007), and corrections for galactic dust extinction were applied. We did not warp the spectral template to match the observed colour.

For the SNe PTF09cl, PTF09hmv, PTF09mwb, PTF09nl, PTF09cne, PTF10uj, PTF10wmm and PTF01xjt, the first observing epoch was before NIR maximum. For the remaining SNe, the first observing epoch was near NIR maximum. When fitting light curves for the SNe, we included the date of \( B \)-band maximum as an extra prior. The \( B \)-band maximum was determined using SITFO (Conley et al. 2008), as was the \( B \)-band stretch (Maguire et al., in preparation).

The \( B \)-band stretches range from 0.8 to 1.15, so there are no fast decliners in our sample. Nor are there any heavily reddened SNe. Furthermore, all of the SNe appear to be spectroscopically normal. The absence of fast decliners and heavily reddened SNe is a potential limitation of the current sample.

The photometry for all SNe is presented in Table 1.

3 TEMPLATE FITTING

To find the light-curve maxima, we fit the Flexible Light-curve InfraRed Template (FLIRT) presented in Mandel et al. (2009) to our \( K \)-corrected photometry4. As we did not observe all SNe Ia before maximum light, we fit the template to the data with an additional constraint. Concurrent observations of SNe Ia at NIR and optical wavelengths provide evidence that NIR maxima occur \( \sim 5 \) d before \( B \)-band maximum (Meikle 2000; Krisciunas et al. 2004). We use this prior assumption as a constraint on the date of NIR maximum, \( t_p \), with a 1\( \sigma \) error of \( \sigma_p = 1 \) d, imposed on the likelihood function. By fitting the light-curve template to the SNe Ia observed before maximum without a prior, we determine that the \( J \)-band maximum occurs 5.36 d before the \( B \)-band maximum with a standard deviation of 0.74 d, and the \( H \)-band maximum occurs 4.28 d before with a standard deviation of 0.70 d. Fig. 1 display the template-fitted light curve for PTF10cne.

The errors in the peak apparent magnitudes in both the \( J \) and \( H \) bands were calculated via 100 Monte Carlo simulations of template fitting to the data, including repeated Gaussian-distributed random samplings within the errors of each point.

To convert apparent magnitudes to absolute magnitudes, we determine the distance modulus for each SN, calculated using the standard flat cosmology of \( (H_0, \Omega_m, \Omega_L) = (70, 0.27, 0.73) \). The redshift measurement for PTF10uj is uncertain since the redshift is measured from the SN and not the host. In this case, we have assumed an uncertainty of \( \sigma_z = 0.005 \) (Maguire et al., in preparation). For all other redshifts, the errors make negligible difference to the distance modulus.

4 RESULTS

The peak apparent magnitudes for each SN Ia in the \( J \) and \( H \) bands are quoted in Table 2. In the \( H \) band, the median absolute

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1 We use the rms about the mean as a measure of the scatter.
2 http://www.eso.org/sci/facilities/paranal/instruments/hawki/ and http://www.gemini.edu/sciops/instruments/niri/
3 IRAF is distributed by the National Optical Astronomy Observatories which are operated by the Association of Universities for Research in Astronomy, Inc. under the cooperative agreement with the National Science Foundation.
4 Photometry to be published in Barone-Nugent et al. (in preparation).
Table 1. The $J$- and $H$-band photometry for each epoch of each SN in our sample. Apparent magnitudes have been corrected for Galactic dust extinction but not $K$-corrected.

| SN name     | MJD   | $m_J$   | $m_H$   |
|-------------|-------|---------|---------|
| PTF09dgc    | 55068.14 | 19.803 ± 0.023 | 18.813 ± 0.021 |
| PTF07gj     | 55939.36 | 19.679 ± 0.073 | 18.770 ± 0.044 |
| PTF08uq     | 55389.35 | 19.815 ± 0.059 | 17.891 ± 0.057 |
| PTF10gxy    | 55422.63 | 19.652 ± 0.069 | 17.790 ± 0.146 |
| PTF10dgc    | 55391.43 | 19.804 ± 0.062 | 17.849 ± 0.079 |
| PTF10fvc    | 55389.35 | 19.810 ± 0.088 | 17.900 ± 0.099 |
| PTF10fbc    | 55389.35 | 19.833 ± 0.092 | 17.908 ± 0.107 |
| PTF10fbc    | 55392.35 | 19.496 ± 0.101 | 17.919 ± 0.118 |

In the $J$ band, the variations between samples are larger and more difficult to understand. We find a median magnitude of $M_J = -18.39 ± 0.06$, compared to median magnitudes of $M_J = -18.48 ± 0.05, -18.29 ± 0.10$ and $-18.73 ± 0.05$ for the SNe in Folatelli et al. (2010), Wood-Vasey et al. (2008) and Krisciunas et al. (2004), respectively. The errors were calculated as $\sigma = 0.17$ mag. and $\sigma = 0.016$ mag and $\sigma = 0.116 ± 0.027$ mag represents a considerable variation in previous samples, including Wood-Vasey et al. (2008) who observed $\sigma_M = 0.28$ mag and $\sigma = 0.116 ± 0.027$ mag. The values for $\sigma_J$ and $\sigma_H$ presented here are the scatter in the magnitude at $J_{max}$ and $H_{max}$, respectively.

Photometric errors make little contribution to the scatter we measure. Peculiar velocities only make a significant contribution if they are large (i.e. $\sim 300 \text{ km s}^{-1}$). The intrinsic magnitudes are assumed to be $18.18 ± 0.01$ in the $J$ band and $18.28 ± 0.015$ in the $H$ band. For peculiar velocities of $150 \text{ km s}^{-1}$, we obtain intrinsic magnitudes of $\sigma = 0.109$ in the $J$ band and $\sigma = 0.077$ in the $H$ band. Part of the scatter we observe may also be due to spectral variability affecting the $K$-corrections.

We note that the light curves of some of the SNe in our sample are not adequately described by the light-curve template that we have used. This is because there is a diversity in light-curve shapes, which may lead to opportunities to reduce the scatter about the Hubble line further. We investigated this by recomputing the scatter after removing SNe with poor light-curve fits. If we eliminate light curves with a $\chi^2$ of freedom greater than 5 (a chance probability of $Q = 0.002$), then we must correct PTF09duc and PTF09dmc in the $J$ band and PTF09dmc, PTF10hmv and PTF10mwb in the $H$ band. The resulting values of the scatter in the $J$ and $H$ bands are 0.139 and 0.097, respectively. The scatter in the $H$ band increases marginally but is noticeably lower for the $J$ band. PTF09duc, PTF10hmv and PTF10mwb exhibit no spectral or $B$-band light-curve peculiarities. However, applying such cuts for the purposes of reducing the scatter needs to be done with care since they may introduce biases that may affect the accuracy with which cosmological parameters can be measured.

The relationship between the peak magnitude in both filters and pseudo-colour ($m_{J,\text{peak}} - m_{H,\text{peak}}$) was also considered. In the optical, there is a strong correlation between colour and luminosity (Tripp 1998; Jha et al. 2007). However, we find no obvious correlation between the $H$-band residual and $I_{\text{max}} - H_{\text{max}}$ pseudo-colour in our sample. On the other hand, we do find a correlation between the $J$-band residual and pseudo-colour. From the least-squares residuals in the y-direction, we find a relation for the line of best fit, where the errors on the coefficients are $1\sigma$ dispersions calculated via
1000 Monte Carlo simulations by randomly varying the points within their error bars:

\[ J - \text{band residual} = (0.91 \pm 0.2) \times (M_{J,\text{max}} - M_{H,\text{max}}) + (0.05 \pm 0.02) \]  

with a scatter about the line of best fit of \( \sigma = 0.071 \).

We also considered the relationship between the peak magnitude in both filters with the angular distance between the SN and the centre of the host galaxy and B-band stretch. However, we find no correlations between any of these quantities. It should be noted that the B-band stretch covered by this sample is limited.

### 4.1 Template fitting with a subset of points

The previous discussion refers to SNe Ia light curves consisting of four data points, which we used to fit the light-curve templates. In this section we discuss the optimal number of epochs that should be observed in the NIR given finite telescope time.

We have already calculated the peak magnitude, \( M_{J,\text{max}} \), for each SN Ia using all data available. We now calculate the peak magnitudes using a subset of the data points of size \( x \), denoted by \( M_{x,\text{max}} \), and find the scatter of these magnitudes around \( M_{x,\text{max}} \). The total scatter in the sample using \( x \) points is denoted by \( \sigma_{J,\text{max,tot}} \). The results for these quantities based on our SN Ia sample are presented in Table 3.

We find in all cases that there is approximately 100th of a magnitude difference in the mean peak magnitude or less between different numbers of epochs used. However, there is additional scatter around the mean which decreases as more points are added. This illustrates the exchange between amount of telescope time used and dispersion measured. Table 3 suggests that the increased accuracy from three points to four is \( \sim 0.01 \text{ mag} \) or less in both the \( J \) and \( H \) bands. Furthermore, using just one point with a known date of maximum can be used while only sacrificing \( \sim 0.04 \text{ mag} \) of accuracy (i.e. 1\sigma scatter) in both the \( J \) and \( H \) bands. Importantly, using just one point in the \( J \) band gives a dispersion that is similar to the best standardized SNe Ia observed in the optical, and using just one point in the \( H \) band can further improve the accuracy of SNe Ia observed in the optical. We also find that, within the range of epochs covered by our data (approximately 10 d before to 15 d after B-band maximum), the amount of additional scatter introduced due to using only one point has no dependence on when this point is taken with respect to B-band maximum. Importantly, single observations after B-band peak are equally as effective as points prior to the peak. We may now compare the standard error of the mean (SEM) for each subset of points. The SEM for a sample of \( n \) points is given by

\[ \text{SEM} = \frac{\sigma}{\sqrt{n}} \]
**NIR Type Ia SNe**

Figure 2. The upper panels show the Hubble diagrams of our sample of 12 SNe (green circles) in the $H$ band (left) and the $J$ band (right), including the Wood-Vasey et al. (2008) sample, the Foltelli et al. (2010) sample and the Krisciunas et al. (2004) sample (blue crosses) for comparison. The red line represents the apparent magnitude that would be observed assuming constant absolute magnitudes of $-18.36$ ($H$ band) and $-18.39$ ($J$ band). The lower panels show the Hubble residual, i.e. the deviation from the red line. The shaded red region, $z < 0.03$, is the region excluded in our sample due to the associated peculiar velocity errors. The solid black lines represent the change in the distance modulus due to a peculiar velocity of $\pm 300$ km s$^{-1}$.

Table 3. The mean peak magnitude and dispersion in the $J$ and $H$ bands when using light curves with $x$ points.

| $x$ | $M_{J, x}-pt$ - $M_J$ | $\sigma_{J, tot}$ | $M_{H, x}-pt$ - $M_H$ | $\sigma_{H, tot}$ |
|-----|-----------------|-----------------|-----------------|-----------------|
| 1   | 0.002           | 0.146           | 0.007           | 0.116           |
| 2   | 0.012           | 0.126           | 0.002           | 0.096           |
| 3   | 0.008           | 0.120           | 0.002           | 0.089           |
| 4   | 0.000           | 0.116           | 0.000           | 0.085           |

Table 4. The SEM for an example of 12 observing nights distributed among three, four, six and 12 SNe.

| $x$ | Sample size | $\sigma_H$ | SEM$_H$ | $\sigma_J$ | SEM$_J$ |
|-----|-------------|------------|---------|------------|---------|
| 1   | 12          | 0.116      | 0.033   | 0.146      | 0.042   |
| 2   | 6           | 0.096      | 0.039   | 0.126      | 0.051   |
| 3   | 4           | 0.089      | 0.045   | 0.120      | 0.060   |
| 4   | 3           | 0.085      | 0.049   | 0.116      | 0.067   |

where $\sigma$ is the standard deviation of the sample. As an illustrative example, we compare observing three SNe over four epochs, four SNe over three epochs, six SNe over two epochs and 12 SNe for a single epoch. In all cases, we assume that we have well-sampled optical light curves for each SN Ia. Each of these would require equal allocation of telescope time. The relative SEM for each method is summarized in Table 4. We find that observing four SNe with only a single epoch delivers a lower SEM than observing fewer SNe over more than one epoch. The larger sample size more than compensates for the higher uncertainty in the peak magnitude.

5 CONCLUSION

We have presented NIR light curves of 12 SNe Ia that are in the Hubble flow. We find that the intrinsic scatter in peak luminosities of SNe Ia in the NIR $J$ and $H$ bands are smaller than previously thought. This is the first sample to display an $H$-band rms scatter as small as $\sigma_H = 0.085 \pm 0.016$ (with a median peak magnitude of $M_H = -18.36$). Our observed $J$-band rms scatter of $\sigma_J = 0.116 \pm 0.027$ (with a median peak magnitude of $M_J = -18.39$) is smaller than reported elsewhere. These results provide distance errors of $\sim 4$ per cent using $H$-band SNe, making them the most precise standard candles for cosmology.

We have also shown that if concurrent optical observations are made, we may use a predicted date of NIR maximum as a constraint when fitting the light curve. With this constraint, we may use as few as one NIR observation within $\sim 5$–10 d of NIR maximum per SN while still achieving scatters of $\sigma_H = 0.116$ and $\sigma_J = 0.146$. As surveys improve over the coming years and more SNe Ia are discovered, single-night NIR observations undertaken concurrently with optical observations will be the most efficient and accurate way to construct samples.

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REFERENCES

Alard C., 2000, A&AS, 144, 363
Brorson T. et al., 2007, in di Salvo T., Israel G. L., Piersant L., Burderi
L., Matt G., Tornambe A., Menna M. T., eds, AIP Conf. Ser. Vol. 924,
Quantitative Spectroscopy of Distant Type Ia Supernovae. American
Institute of Physics, Maryland, 415
Casali M. et al., 2006, Proc. SPIE, 6269, 62690W
Conley A. et al., 2008, ApJ, 681, 482
Conley A. et al., 2011, ApJS, 192, 1
Folatelli G. et al., 2010, ApJ, 139, 120
Hamuy M. et al., 1996, AJ, 112, 2408
Hodapp K., Hora J., Graves E., Irwin E., Yamada H., Douglass J., Young T.,
Robertson L., 2000, Proc. SPIE, 4008, 1334
Hsiao E., Conley A., Howell D., Sullivan M., Pritchett C., Carlberg R.,
Nugent P., Phillips M., 2007, ApJ, 663, 1187
Jha S., Riess A., Kirshner R., 2007, ApJ, 659, 122
Kasen D., 2006, ApJ, 649, 939
Kattner S. et al., 2012, PASP, 124, 127
Krichnan K., Phillips M., Suntzeff N., 2004, ApJ, 602, L81
Law N. et al., 2009, PASP, 121, 1395
Mandel K., Narayan G., Kirshner R., 2011, ApJ, 731, 120
Mandel K., Wood-Vasey W., Friedman A., Kirshner R., 2009, ApJ, 704,
629
Meikle W., 2000, MNRAS, 314, 782
Persson S., Murphy D., Krzeminski W., Roth M., Rieke M., 1998, ApJ, 116,
2475
Phillips M., 1993, ApJ, 413, L105
Phillips M., 2011, preprint (arXiv:1111.4463)
Rau A. et al., 2009, PASP, 121, 1334
Silverman J., Ganeshalingam M., Li W., Filippenko A., Alexei V., 2012,
MNRAS, in press
Smith A. et al., 2011, MNRAS, 412, 1309
Sullivan M. et al., 2010, MNRAS, 406, 782
Tripp R., 1998, A&A, 331, 815
Walker E. et al., 2011, MNRAS, 410, 1262
Wood-Vasey W. et al., 2008, ApJ, 689, 377

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