Discovery of PSR J0523-7125 as a Circularly Polarized Variable Radio Source in the Large Magellanic Cloud

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Abstract

We report the discovery of a highly circularly polarized, variable, steep-spectrum pulsar in the Australian Square Kilometre Array Pathfinder (ASKAP) Variables and Slow Transients (VAST) survey. The pulsar is located about 1° from the center of the Large Magellanic Cloud, and has a significant fractional circular polarization of ∼20%. We discovered pulsations with a period of 322.5 ms, dispersion measure (DM) of 157.5 pc cm−3, and rotation measure (RM) of +456 rad m−2 using observations from the MeerKAT and the Parkes telescopes. This DM firmly places the source, PSR J0523−7125, in the Large Magellanic Cloud (LMC). This RM is extreme compared to other pulsars in the LMC (more than twice that of the largest previously reported one). The average flux density of ∼1 mJy at 1400 MHz and ∼25 mJy at 400 MHz places it among the most luminous radio pulsars known. It likely evaded previous discovery because of its very steep radio spectrum (spectral index α ≈ −3, where Sν ∝ ν−α) and broad pulse profile (duty cycle ≥35%). We discuss implications for searches for unusual radio sources in continuum images, as well as extragalactic pulsars in the Magellanic Clouds and beyond. Our result highlighted the possibility of identifying pulsars, especially extreme pulsars, from radio continuum images. Future large-scale radio surveys will give us an unprecedented opportunity to discover more pulsars and potentially the most distant pulsars beyond the Magellanic Clouds.

Unified Astronomy Thesaurus concepts: Pulsars (1306); Neutron stars (1108); Radio transient sources (2008); Radio continuum emission (1340)

1. Introduction

Since the discovery of the first pulsar (Hewish et al. 1968), a lot of effort has been devoted to developing efficient and sensitive search algorithms (e.g., Ransom 2001; Ransom et al. 2003). The traditional radio pulsar search procedures focus on the Fourier domain or time domain to identify periodic signals (see Chapter 6 in Lorimer & Kramer 2012). While this has been fruitful and has discovered more than ∼3000 pulsars (Manchester et al. 2005), abnormal pulsars such as short orbital period binary systems or strongly scattered objects are more difficult to detect. Indeed, some algorithms have been developed (e.g., for acceleration and “jerk” searches; Ransom 2001; Andersen & Ransom 2018) to explore a broader parameter space, but they are computationally expensive for untargeted sky surveys.

Continuum images have long been considered an effective way to select pulsar candidates based on their steep spectral indices (e.g., Backer et al. 1982; Damico et al. 1985; Strom 1987; Frail et al. 2016b; Bhakta et al. 2017; Maan et al. 2018). Since continuum images are sensitive to pulsar
Pulsars are known to be compact enough to exhibit variability due to interstellar scintillation, caused by irregularities in the turbulent interstellar medium (Rickett 1990). Dai et al. (2016) proposed a new detection technique based on this, using diffractive interstellar scintillation in variance images. The scintillation behavior, plus potential intrinsic fluctuations (e.g., intermittent behavior; Lyne 2009), will cause strong flux density variations, and therefore pulsars can sometimes be detected in general radio variability surveys (e.g., Bell et al. 2016; Murphy et al. 2021).

Another approach for identifying pulsars in continuum surveys is through circularly polarized emission (Gaensler et al. 1998; Kaplan et al. 2019). Only a few types of sources are known to be more than a few percent circularly polarized, and they are usually pulsars (e.g., Dai et al. 2015; Lenc et al. 2018) or stellar objects (e.g., Pritchard et al. 2021). Thus, highly circularly polarized sources that lack a deep optical/infrared counterpart are strong pulsar candidates. To date, there have only been two large-scale circular polarization surveys. Lenc et al. (2018) conducted the first all-sky circular polarization survey using the Murchison Widefield Array (MWA) at 200 MHz, and detected 14 known pulsars in the untargeted survey. After that, Pritchard et al. (2021) performed a circular polarization survey for radio stars with the Australian Square Kilometre Array Pathfinder (ASKAP; Hotan et al. 2021) at 888 MHz, and identified 33 known pulsars.

We use the data from the ASKAP Phase I Pilot survey for Variables and Slow Transients (VAST-P1; Murphy et al. 2021) processed with the VAST pipeline (Pintaldi et al. 2021) to search for variable and transient sources in the two ASKAP fields covering the Magellanic Clouds. Here we report the discovery of a highly variable, circularly polarized, steep-spectrum source VAST J052348.6−712552 in this continuum survey. After targeted MeerKAT and Parkes follow-up observations, we identified it as a new pulsar, PSR J0523−7125 located in the Large Magellanic Cloud (LMC). The extragalactic distance makes it one of the most luminous pulsars known at both 400 MHz (\(L_{400} \approx 6.5 \times 10^7\) mJy kpc\(^2\)) and 1400 MHz (\(L_{1400} \approx 2.5 \times 10^8\) mJy kpc\(^2\)). We present our observations and results in Section 2. In Section 3, we discuss the nature of the pulsar and prospects for identifying future pulsars through continuum imaging surveys.

2. Observations and Results

2.1. ASKAP Discovery

We observed two 30 deg\(^2\) fields (field names VAST 0530−68A and VAST 0127−73A) centered on the Magellanic Clouds, six times between 2019 August and 2020 January. Each observation had an integration time of 12 minutes, achieving a typical rms sensitivity of 0.25 mJy beam\(^−1\) and angular resolution of 12″ at a central frequency of 888 MHz with a bandwidth of 288 MHz. VAST-P1 incorporated data from the Rapid ASKAP Continuum Survey (RACS; McConnell et al. 2020) as the first epoch, which has the same observing frequency, but a longer integration time of 15 minutes, resulting in an rms sensitivity of 0.20 mJy beam\(^−1\) for regions near the Magellanic Clouds. Details of data reduction for these surveys are given in McConnell et al. (2020) and Murphy et al. (2021).

We conducted a search for highly variable sources in the two fields using the VAST transient detection pipeline (Pintaldi et al. 2021; Murphy et al. 2021). We selected candidates with a variability index \(\mathcal{V} > 1.0\sigma_{\mathcal{V}}\) (equivalent to the fractional variability used by other surveys) and reduced chi-square \(\eta > 2.0\sigma_{\eta}\) (equating to values of \(\mathcal{V} > 0.295\) and \(\eta > 6.479\)), where \(\sigma\) is the standard deviation measured by fitting a Gaussian function to the distributions of both metrics in logarithmic space (following the calculations in Rowlinson et al. 2019). We identified ∼27 candidates as highly variable, compact sources. Two of them were detected with strong circular polarization, and VAST J052348.6−712552 was the only one lacking a clear multiwavelength association/identification (the other one is associated with a stellar object). The
fractional circular polarization of VAST J052348.6–712552 ($f_p = |S_V|/S_I$) is about 15%–30% (see the cutout in Figure 1). No linear polarization was detected for this source in VAST-P1, placing a $3\sigma$ upper limit of $\lesssim 0.9$ mJy beam$^{-1}$ ($\lesssim 10\%$ fractional polarization) on the total linear polarization ($P = \sqrt{Q^2 + U^2}$). The radio observations are summarized in Table 1 and shown in Figure 2. Other variable sources and the general survey results will be published in a subsequent paper.

2.2. MeerKAT and Parkes Observations

We observed VAST J052348.6–712552 with the MeerKAT radio telescope (pulsar search mode and continuum imaging mode simultaneously) for 2.5 hr at a central frequency of 1284 MHz (bandwidth of 856 MHz) on 2021 August 25. We searched the MeerKAT (pulsar mode) data for pulsar candidates using the standard Fourier domain search procedure. This was done using PULSAR\_MINER\(^{23}\) (see Ridolfi et al. 2021 for more details), an automated pipeline based on the PRESTO\(^{24}\) pulsar searching package (Ransom 2011). After cleaning the observing data by removing the frequency channels affected by strong radio frequency interference (RFI), we generated dispersed time series, correcting for the interstellar dispersion.

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Note. Archival surveys giving best constraints are selected for optical/infrared. In the column of flux density, radio continuum observation refers to the total intensity (Stokes I) with units of mJy beam$^{-1}$, optical/infrared observation refers to the magnitude (Vega system), and X-ray observation refers to the unabsorbed flux (0.2–10 keV). Flux densities are not available for Parkes pulsar observations. No Stokes V images/catalogs are available for SUMSS, ASKAP SB8178, or ASKAP SB8532. The nondetection upper limits of radio flux measurements are at $3\sigma$ confidence level. For the three Stokes V nondetections in the VAST-P1, we combined these three epochs and measured a $3\sigma$ peak at the position of VAST J052348.6–712552 from the combined mean image, which yields a $|S_V|/S_I \approx 15\%$, consistent with other detections. The upper limits of other multiwavelength observations are listed at the $3\sigma$ confidence level. Due to bad image quality for the ATCA observation at 9000 MHz on 2020 August 12 (compact array configuration) and for the ATCA stokes V image on 2021 September 10 (strong artifacts); those images are discarded and not reported here.

References. (1) Mauch et al. (2003); (2) Hurley-Walker et al. (2017); (3) Pennock et al. (2021); (4) McConnell et al. (2020); (5) Zaritsky et al. (2004); (6) Cioni et al. (2011).

| Obs. Date (UTC) | Telescope | Duration | Band (MHz) | Flux Density (mJy beam$^{-1}$) | $|S_V|/S_I$ | Survey/Project ID | Ref. |
|----------------|-----------|----------|------------|-------------------------------|------------|--------------------|------|
| 1997 Dec 22    | MOST      | 11.5 hr  | 843        | 8.6 ± 0.9                     | ...        | ...                | SUMSS 1 |
| 2013 Nov 15    | MWA       | ~8 minutes | 72–231   | $\lesssim 7$ (3$\sigma$) | ...        | ...                | GLEAM 2 |
| 2019 Mar 13    | ASKAP     | 12 hr 40 minutes | 1420   | 3.6 ± 0.1                     | ...        | ...                | SB8178 |
| 2019 Apr 20    | ASKAP     | 12 hr 40 minutes | 888    | 4.3 ± 0.1                     | ...        | EMU/SB8532         | 3    |
| 2019 May 7     | ASKAP     | 15 minutes | 888     | 2.07 ± 0.15                   | 1.00       | <48%               | RACS 4 |
| 2019 Aug 28    | ASKAP     | 12 minutes | 888     | 2.44 ± 0.31                   | <1.24      | <51%               | VAST-P1 |
| 2019 Dec 19    | ASKAP     | 12 minutes | 888     | 7.24 ± 0.25                   | 1.38 ± 0.22 | 19%               | VAST-P1 |
| 2020 Jan 10    | ASKAP     | 12 minutes | 888     | 4.73 ± 0.25                   | <1.04      | <22%               | VAST-P1 |
| 2020 Jan 24    | ASKAP     | 12 minutes | 888     | 5.99 ± 0.22                   | 1.48 ± 0.21 | 25%               | VAST-P1 |
| 2020 Jan 25    | ASKAP     | 12 minutes | 888     | 5.13 ± 0.21                   | 1.20 ± 0.20 | 23%               | VAST-P1 |
| 2020 Apr 13    | Parkes    | 30 minutes | 704–4032 | ...                           | ...        | ...                | P1069 |
| 2020 Jul 1     | Parkes    | 60 minutes | 704–4032 | ...                           | ...        | ...                | P1069 |
| 2020 Aug 12    | ATCA      | 2 hr      | 2100     | 0.836 ± 0.036                 | <0.159     | <19%               | C3363, NAPA |
| 2020 Aug 12    | ATCA      | 2 hr      | 5500     | <0.122                        | ...        | C3363, NAPA        |      |
| 2020 Aug 12    | ATCA      | 2 hr      | 9000     | ...                           | ...        | C3363, NAPA        |      |
| 2020 Aug 29    | ASKAP     | 12 minutes | 888     | 3.89 ± 0.24                   | <1.06      | <27%               | VAST-P1 |
| 2020 Nov 18    | Parkes    | 135 minutes | 704–4032 | ...                           | ...        | ...                | P1069 |
| 2021 Apr 1     | ASKAP     | 12 minutes | 888     | 7.40 ± 0.23                   | 1.47 ± 0.19 | 20%               | VAST-P2-low |
| 2021 Apr 2     | ASKAP     | 12 minutes | 1367.5   | 1.47 ± 0.26                   | <1.31      | <89%               | VAST-P2-mid |
| 2021 Apr 25    | ATCA      | 80 minutes | 2100     | 0.687 ± 0.038                 | <0.148     | <22%               | C3431 |
| 2021 Apr 25    | ASKAP     | 60 minutes | 5500     | <0.128                        | ...        | C3431 |
| 2021 Apr 25    | ATCA      | 60 minutes | 9000     | <0.083                        | ...        | C3431 |
| 2021 Jun 4     | ATCA      | 4 hr      | 2100     | 0.364 ± 0.020                 | 0.076 ± 0.013 | 21%            | C3431 |
| 2021 Jun 20    | ASKAP     | 13 minutes | 888     | 7.6 ± 0.17                    | 1.5 ± 0.18  | 20%               | VAST-P2-low |
| 2021 Aug 25    | MeerKAT   | 2.5 hr    | 1284     | 0.860 ± 0.007                 | 0.137 ± 0.005 | 16%            | DDT-20210818-TM-01 |
| 2021 Sep 10    | ATCA      | 5 hr      | 2100     | 0.176 ± 0.018                 | ...        | ...                | C3431 |

1995–1999 LCST/GCC ... $U$ $>20.7^m$ MCPS 5
... $B$ $>22.6^m$ ...
... $V$ $>22.5^m$ ...
... $I$ $>21.2^m$ ...

2009–2012 VISTA 7 200 s $Y$ $>21.1^m$ VMC 6
7 200 s $J$ $>20.9^m$ ...
27 000 s $K_s$ $>20.4^m$ ...

2021 May 29 Swift/XRT 2 250 s 0.2–10 keV $<1.3 \times 10^{-13}$ Target ID: 14338
(erg s$^{-1}$ cm$^{-2}$)
with dispersion measure (DM) trial values in the range 2.0–300 pc cm$^{-3}$, with steps of 0.05 pc cm$^{-3}$. Each time series was Fourier transformed and the resulting power spectrum searched for prominent periodicities using PRESTO’s accelersearch routine. The latter is sensitive to both isolated and binary pulsars, by accounting for possible Doppler shifts of the pulsar spin frequency in the Fourier domain due to orbital motion (see Ransom et al. 2002 for a detailed discussion of the acceleration search technique). For our search, we allowed a maximum drift of up to 200 Fourier bins, using the $z$max 200 option of accelersearch. We identified a strong pulsar candidate, PSR J0523$-$7125, with a period of 322.5 ms at a DM of 157.5 pc cm$^{-3}$. No significant acceleration was detected within the 2.5 hr of the observation, indicating that the candidate pulsar is likely isolated, although further timing is ongoing. Figure 3 shows the initial discovery plot of the pulsar. It has a very steep spectrum, consistent with the measurement in continuum imaging (see described below) and a wide pulse profile with duty cycle $\sim$35% (using the pulse width at 50% maximum, $W_{50}$). Note the duty cycle may be underestimated as it is difficult to identify the off-pulse baseline for a (wide-pulsed) pulsar like this. We also measured a RM = +456 ± 6 rad m$^{-2}$ using the rmpfit tool in PSRCHe (Hotan et al. 2004) and found a significant linear polarization ($\geq$50%) after RM correction (see pulse profiles in Figure 4). Further flux calibration is necessary to measure precise polarization properties. This pulsar has not been recorded in any known pulsar catalog (Manchester et al. 2005) or in the Pulsar Survey Scraper, which records newly discovered pulsars prior to publication.

We also processed the MeerKAT (simultaneous) continuum data using OXKAT (Heywood 2020), which uses: CASA for basic flagging, cross calibration, and splitting out measurement sets; TRICOLOUR for further flagging; CUBICAL (Kenyon et al. 2018) for self-calibration; and WSCLEAN (Offringa et al. 2014) for continuum imaging. We used PMNJ0408$-$6545 for bandpass and flux calibration, and PMNJ0420$-$6223 for phase calibration. We detected VAST J052348.6$-$712552 with a total flux density (Stokes I) of 860 $\pm$ 7 Jy beam$^{-1}$ (at $\geq$100$\sigma$ confidence level), and circularly polarized flux density (Stokes V) of 137 $\pm$ 5 Jy beam$^{-1}$. The source is unsolved in the image, which has an angular resolution of 7.5. In contrast to the pulsar mode result, no linear polarization was detected above 3$\sigma$ threshold, and a rotation measure (RM) synthesis found no significant polarized intensity in $|\text{RM}| < 1200$

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25 https://pulsar.cgca-hub.org

26 https://github.com/ska-sa/tricolour
We will discuss this inconsistency in Section 3.1. We obtained the (Stokes I) spectrum from 8 sub-band images and noticed an unusual upturn at higher frequency (see Figure 5). We fitted a smoothly broken power law to the spectrum, and found an upturn at $1.36 \pm 0.02$ GHz with a spectral index $\alpha = -4.40 \pm 0.11$ in the lower frequency range and $\alpha = +2.34 \pm 0.51$ in the upper frequency range. We also obtained lightcurves (with 10 minutes, 2 minutes, and 8 s resolution, respectively) within the 2.5 hr observation, and found no significant variability on those timescales.

We observed VAST J052348.6−712552 with the 64 m Parkes telescope on 2020 April 13, 2020 July 1, and 2020 November 18 using the pulsar searching mode with the Ultra-Wideband Low (UWL) receiver (which provides simultaneous frequency coverage from 704 to 4032 MHz; Hobbs et al. 2020). We used PRESTO (Ransom 2001) to perform a standard pulsar search of DMs spanning 0−300 pc cm$^{-3}$, plus an acceleration search up to $\sim 200$ m s$^{-2}$. Our initial search of the Parkes observations found some candidate pulsars, but none were convincing. After detecting the pulsar with MeerKAT we reinspected our data and found a pulsar candidate with a similar period and DM. It had not been identified as a convincing candidate in our initial search since the signal is very weak (near our searching threshold SNR of $\sim 8$), possibly due to the wide pulse profile and the steep spectrum (so it is only bright at low frequencies). We then reanalyzed the Parkes observations using the low-frequency data only, and successfully detected this pulsar at frequencies $\sim 2$ GHz. Our Parkes data also measured a similar RM = $+457.6 \pm 0.1$ rad m$^{-2}$ using RMFIT.

We note that this sky region was searched as part of the survey conducted by Ridley et al. (2013) using the Parkes telescope, who claimed a flux density limit of 0.05 mJy at 1400 MHz, well below the flux density of PSR J0523−7125. However, we were able to identify a weak candidate in their data as well, with parameters consistent with PSR J0523−7125. Note that the Ridley et al. (2013) limit assumed a 5% pulse width (compared to 35% for PSR J0523−7125) and assumed that sources were at the centers of the beams of the Parkes Multi-Beam Receiver. Accounting for our pulse width increases the limit by $\sim 3$ while accounting for the position (as the pulsar is close to the half-power point of one of the beams) increases it by a further factor of $\sim 2$. This gives a limiting flux density of about 0.3 mJy, still less than what we observe, but the remaining differences may be due to a combination of scintillation, lost bandwidth due to radio frequency interference, and similar effects.

57 Following https://docs.astropy.org/en/stable/api/astropy.modeling.powerlaws.SmoothlyBrokenPowerLaw1D.html.
2.3. ATCA and Archival Radio Observations

Prior to having identified the pulsations in PSR J0523−7125, we carried out follow-up observations of VAST J052348.6−712552 with the Australia Telescope Compact Array (ATCA) on 2020 August 12, 2021 April 25, 2021 June 4, and 2021 September 10 respectively, with a 2 GHz bandwidth centered at 2.1 GHz, 5.5 GHz, and 9.0 GHz (project code C3363, PI: Murphy; and project code C3431, PI: Pritchard). We reduced the data using MIRIAD (Sault et al. 1995) with PKS B1934−638 and B0530−727 as the flux and phase calibrators, respectively, and imaged the data using CASA (McMullin et al. 2007). We detected the source with a flux density of 0.3−0.9 mJy at 2.1 GHz, but did not detect it at 5.5 GHz or 9.0 GHz in any epoch. We measured a circular polarization flux density of 0.076 ± 0.013 mJy beam−1 at 2.1 GHz for the observation on 2021 June 4, implying a polarization fraction of ∼20%. The source is unresolved in all detected ATCA images, with an angular resolution of 3−6″. We measured the

Figure 4. Folded pulse profiles of PSR J0523−7125, from the MeerKAT observation (left) and the Parkes observation (right). We show the pulse intensity as a function of frequency over the full bandpass for MeerKAT (lower left) and only the lower UWL bandpass (lower right). The integrated pulse profiles (top) are summed over frequencies ≤1100 MHz, where the signal-to-noise ratio is best. We also show the polarized intensity profiles for both MeerKAT and Parkes observation (with RM correction). Note the linear polarization (P) is (unphysically) higher than total intensity (I) at some pulse phases, possibly due to a combination of RFI issues and “off-pulse baseline” subtraction uncertainty.

Figure 5. Left: the spectrum measured at the MeerKAT follow-up, with Stokes I (total intensity) in the upper panel and fractional circular polarization (|SV|/SI) in the lower panel. The error bars are smaller than the marker size. We split the 856 MHz bandwidth into 8 sub-bands, and fitted with a smoothly broken power law. Right: the spectra measured at different ATCA follow-up epochs, with error bars smaller than the marker size. We equally split the 2 GHz bandwidth into 4 sub-bands (centered at 1332 MHz, 1844 MHz, 2356 MHz, and 2868 MHz respectively). The open markers represent 5σ upper limits. We fit the spectra with a single power law in each case, and the fitted spectral indices are shown in the legend. See Sections 2.3 and 2.2 for details.

28 The array was in a compact configuration EW352 on 2020 August 12, and only baselines including antenna 6 were included, corresponding to a baseline range of 4087−4439 km.
spectral index by splitting the L-band (2.1 GHz) into four sub-bands, and found that it was very steep, varying between \( \alpha \approx -4 \) to \( \sim -2 \). The spectra can be fitted using a single power law, and their distributions are shown in Figure 5. The ATCA observations allowed us to measure a more accurate position compared to ASKAP \( (\alpha = 0^\circ 23'48''66, \delta = -71^\circ 25'52''58 \) in J2000 coordinates), with positional uncertainties of \( \sim 0.15 \) arcsec in each coordinate.

We also detected VAST J052348.6–712552 in archival radio observations including the Sydney University Molonglo Sky Survey (SUMSS; Mauch et al. 2003) and an ASKAP commissioning observation (SB8532; Pennock et al. 2021). We obtained a higher-time resolution lightcurve from the latter (by imaging 9 \times 7 minutes scans separated by 1.2 hr) and found no significant variability within the 12 hr 40 minutes observation. This area is also observed by the GaLactic and Extragalactic All-sky Murchison Widefield Array (GLEAM) survey (Hurley-Walker et al. 2017), and the nondetection yields a 3\( \sigma \) upper limit of \( \sim 40 \) mJy at 72–231 MHz. All of these observations are listed in Table 1.

### 2.4. Multiwavelength Imaging Observations

Prior to initiating our pulsar searches we searched archival multiwavelength data including Gaia (Gaia Collaboration et al. 2018), the Wide-field Infrared Survey Explorer (WISE; Wright et al. 2010), the ROSAT All-Sky Survey (Boller et al. 2016), and the Fermi Large Area Telescope source catalog (Acero et al. 2015). No multiwavelength counterparts were found within a 5\( \prime \) radius in Gaia or WISE, or within a 2\( \prime \) radius in ROSAT or Fermi.

The most sensitive infrared data available for this region is the VISTA Magellanic survey catalog (VMC; Cioni et al. 2011). We found no sources within a 5\( \sigma \) position uncertainty, leading to \( \sigma \) limiting magnitudes \( Y > 21.1 \) mag, \( J > 20.9 \) mag, and \( K_s > 20.4 \) mag in this region. The position is also covered by the Magellanic Clouds Photometric Survey (MCPS; Zaritsky et al. 2004), and its nondetections of the source indicate \( U > 0.7 \) mag, \( B > 22.6 \) mag, \( V > 22.5 \) mag, and \( I > 21.2 \) mag.

We conducted a Neil Gehrels Swift Observatory Target of Opportunity (ToO) observation (target ID: 14338) on 2021 May 29 using the X-ray Telescope (XRT) in the photon counting mode (exposure time of 2250 s). There is one count within a 15 arcsec radius of VAST J052348.6–712552, implying a count rate upper limit of \( \sim 0.0013 \, \text{s}^{-1} \) (0.2–10 keV). We therefore estimated a tentative upper limit to the unabsorbed flux of \( \sim 1.3 \times 10^{-13} \, \text{erg s}^{-1} \text{cm}^{-2} \) (0.2–10 keV) on an HI column density \( N_H = 0.6 \times 10^{22} \, \text{cm}^{-2} \) (similar to the LMC magnetar SGR 0526–66; Park et al. 2012) and a power-law photon index of \( \Gamma = 2 \) using the HEASARC web-based PIMMS. Details of these multiwavelength observations are summarized in Table 1.

### 3. Discussion

VAST J052348.6–712552 was initially selected in the VAST continuum survey for follow-up due to its high variability and strong circular polarization (\( f_p \sim 20\% \)), suggesting it is likely a pulsar or stellar object. The absence of any deep infrared/optical counterpart (3\( \sigma \) limit \( J > 20.9 \) mag) strongly suggested it was a pulsar. Pulsar searches with MeerKAT and Parkes confirmed the pulsar nature of VAST J052348.6–712552. This pulsar, PSR J0523–7125,

has a period of 322.5 ms, DM of 157.5 pc cm\(^{-3}\), and a wide pulse profile. Our ATCA observations show very steep spectral indices ranging from \( \alpha \approx -4 \) to \( \sim -2 \). The MeerKAT spectrum shows an unusual upturn at 1.36 GHz from \( \alpha = -4.4 \) to \( \alpha = 2.34 \). We will discuss the properties of this pulsar in detail in Section 3.1.

#### 3.1. Pulsar Properties

As shown in Figure 1, we know the source is located in the direction of the LMC. The maximum DM contribution from the Milky Way along this line of sight is \( \sim 50–60 \text{ pc cm}^{-3} \) (NE2001 and YMW16; Cordes & Lazio 2002; Yao et al. 2017), and there is no known Galactic HII region along this line of sight. With a DM of 157.5 pc cm\(^{-3}\), PSR J0523–7125 almost certainly resides in the LMC. We also note this DM is consistent with the DMs of other pulsars found in the LMC \( \sim 50–270 \text{ pc cm}^{-3} \), e.g., McConnell et al. 1991; Manchester et al. 2006; Ridley et al. 2013; Johnston et al. 2022.

We can calculate the monochromatic radio luminosity density of PSR J0523–7125 at 400 MHz and 1400 MHz respectively \( (L_d = S_d d^2, \text{where} S_d \text{is the flux density and} d \text{is the pulsar distance}) \). Our ATCA, MeerKAT, and ASKAP midband observations covered 1400 MHz, and the measured average flux density \( S_{1400} \approx 1 \text{ mJy} \). The flux density at 400 MHz is therefore \( S_{400} \approx 25 \text{ mJy} \) assuming a median spectral index \( \alpha \approx -2.6 \) (Figure 5). The flux density at 400 MHz would be ten times higher for the steepest spectral index that we measure \( (\alpha \approx -4.4) \). The pulsar distance \( d \) is assumed to be the LMC distance of 50 kpc (Pietrzyński et al. 2013), which is reasonable since the uncertainty in luminosity should be dominated by the uncertainty in the flux density. We can then estimate \( L_{400} \approx 6.3 \times 10^4 \text{ mJy kpc}^2 \) and \( L_{1400} \approx 2.5 \times 10^3 \text{ mJy kpc}^2 \), which is one of the most luminous pulsars to date (see Figure 6). Only one known pulsar with a well-constrained distance, B1641–45 (Komesaroff et al. 1973; Frail et al. 1991), has a comparable luminosity to PSR J0523–7125 \( \sim 7.6 \times 10^4 \text{ mJy kpc}^2 \) at 400 MHz and \( \sim 9 \times 10^3 \text{ mJy kpc}^2 \) at 1400 MHz. We exclude other high-luminosity pulsars, like PSR J0134–2937, from consideration due to their significant distance uncertainties.

Ridley et al. (2013) calculated the luminosity function of the LMC pulsars and noticed a discrepancy between the Galactic and LMC luminosity distributions. They attributed the discrepancy to a bias due to the small sample size of the LMC pulsars and concluded that the luminosity function for the LMC pulsars is consistent with its counterpart in the Galactic disk. They also suggested that the maximum 1400 MHz radio luminosity for LMC pulsars is approximately 1000 mJy kpc\(^2\) based on their sample. Our pulsar is about a factor of \( \sim 2 \) more luminous than this limit. With the discovery of PSR J0523–7125, we can revisit the luminosity function of the LMC pulsars by fitting the high-luminosity tail of the distribution as a power law, i.e., log \( N = F \log L + G \), where \( F \) is the slope of the distribution (Lorimer et al. 2006). Based on the ATNF pulsar catalog (Manchester et al. 2005), we estimate \( F \approx -1.30 \pm 0.03 \) for Galactic pulsars with luminosities above 30 mJy kpc\(^2\) (consistent with other work, e.g., Lorimer 2004; Ridley et al. 2013), and \( F \approx -1.6 \pm 0.2 \) for LMC pulsars with luminosities above 125 mJy kpc\(^2\). Although the calculation is still limited by small number statistics for the LMC pulsars and the distance uncertainties for Galactic pulsars, our result shows that the LMC pulsars are consistent with the luminosity
function of their Galactic counterparts (as suggested by Manchester et al. 2006; Ridley et al. 2013).

PSR J0523–7125 has a wide pulse profile with \( W_{\text{psr}} \approx 100 \text{ ms at } \sim 1 \text{ GHz. Figure 6 shows the duty cycle distribution of known pulsars, and PSR J0523–7125 generally has a larger duty cycle than others (either Galactic or extragalactic pulsars). The observed pulse width \( W \) can be decomposed as \( W = \sqrt{\tau_{\text{int}}^2 + \tau_{\text{sc}}^2 + \tau_{\text{DM}}^2} \), where \( \tau_{\text{int}} \) is the intrinsic pulse width, \( \tau_{\text{sc}} \) is the pulse broadening due to multipath scattering, and \( \tau_{\text{DM}} \) is the pulse broadening due to dispersion (Lorimer & Kramer 2012). Both \( \tau_{\text{sc}} \) and \( \tau_{\text{DM}} \) are related to free electrons in the interstellar medium (i.e., the DM; Bhat et al. 2004). Following the YMW16 model (which includes the distribution of free electrons in the Galaxy and the Magellanic Clouds), the predicted scattering delay for Magellanic Clouds pulsars along this line of sight is \( \tau_{\text{sc}} \ll 1 \text{ ms at } 1100 \text{ MHz. This suggests that the wide pulse profile is largely intrinsic. As noted in Section 2.2, the pulse width may be further underestimated due to the difficulty of identifying the off-pulse baseline. If this is the case, PSR J0523–7125 could potentially have a nearly 100% duty cycle, i.e., a possible nearly aligned rotator (the magnetic and rotation axis of the star are aligned); e.g., Young et al. 2010). This assumption could also explain why the pulsar is so bright in continuum images, but hard to detect in a pulsar search. The pulsar calibration is ongoing and we will discuss it further in a later paper.

Our MeerKAT pulsar data measured a RM = +456 rad m\(^{-2}\), and the Parkes data confirmed this result. However, we cannot find any significant linear polarization in MeerKAT continuum data in an RM range from \(-1200\) to \(1200\) rad m\(^{-2}\). This discrepancy can be attributed to the potential swing of the polarization position angle in the pulse phase (e.g., following the characteristic S-shaped curve of the rotating vector model; Radhakrishnan & Cooke 1969), making the effective linear polarization fraction very diluted when integrated into the continuum image. Further precise polarimetry can confirm this. We noticed this [RM] is about twice the most extreme [RM] of known LMC pulsars, PSR J0540–6919 with [RM] = −246 rad m\(^{-2}\) (Johnston et al. 2022), which is located in a supernova remnant. We checked the UM/CTIO Magellanic Cloud emission-line survey (MCELS; Smith & MCELS Team 1998) and found there is some faint, diffuse emission around our pulsar in the H\(\alpha\) map (with an angular scale \(\sim 13\) arcmin). We got a H\(\alpha\) intensity \(I_{\text{H\alpha}} \approx 1.3 \times 10^{-16}\) erg cm\(^{-2}\) s\(^{-1}\) arcsec\(^{-2}\) from the MCELS map after extinction correction (following the method in Gaensler et al. 2005), and thus the emission measure (EM) is \(\sim 70\) pc cm\(^{-6}\). The foreground contribution to the DM in this line of sight is about \(60\) pc cm\(^{-6}\) (Yao et al. 2017), and to the RM is about \(+31\) rad m\(^{-2}\) (Mao et al. 2012). After subtracting the Galactic foreground contribution, the estimated parallel magnetic field \(B_{\parallel} \approx 1.23\) RM\(_{\text{LMC}}\)/DM\(_{\text{LMC}}\) \(\approx 5\) \(\mu\)G, and the electron column density \(n_e \approx \text{EM}/\text{DM}_{\text{LMC}} \approx 0.7\) cm\(^{-3}\). The occupation length, \(fL\), of ionized gas is \(\text{DM}_{\text{LMC}}/\text{EM} \approx 150\) pc, where \(L\) is the projected length and \(f\) is the filling factor (the fraction of the line of sight for free electron density). For reasonable values of \(f, L/d\) is of the order of the angular scale of the H\(\alpha\) emission (where \(d \sim 50\) kpc for the LMC). If we assume the angular size is \(13\) arcmin (as we observed from the MCELS image), the inferred filling factor \(f \approx 0.8\) and the length \(L \approx 180\) pc. Those values are reasonable for an evolved HII region. If the pulsar turns out to be young in future timing analysis, it is likely embedded in this diffuse gas region.

In our continuum observations, VAST J052348.6–712552 shows strong variability with a modulation index \(\sim 43\%\) at 888 MHz over timescales of \(\sim\)days (see the lightcurve in Figure 2). There is no significant continuum variability within the 2.5 hr MeerKAT observation or the 12 hr ASKAP EMU
observation. A structure–function analysis shows that the variability timescale is ~17 d, though this is poorly constrained due to the absence of samples between 1 d and 14 d. This strong variability could be intrinsic (e.g., pulsar nulling; Backer 1970) or related to propagation effects (e.g., diffraactive scintillation; Rickett 1990). The diffraactive scintillation band-width follows \( \Delta f_{\text{DISS}} \sim C_1/2\pi f \tau_{\text{sc}} \) MHz (Cordes & Lazio 2002), where \( C_1 = 1.16 \) assuming a Kolmogorov spectrum and \( \tau_{\text{sc}} \) is the scattering time. The predicted scattering delay for the LMC pulsars in this line of sight is \( \tau_{\text{sc}} \sim 10^{-2} \) ms based on the YMW16 model, and thus \( \Delta f_{\text{DISS}} \sim 10^{-2} \) MHz. The scintillation strength, \( u = \sqrt{f/\Delta f_{\text{DISS}}} \) where \( f \) is the observing frequency, is therefore \( u \gg 1 \), implying a strong scattering regime (i.e., diffraactive scintillation and/or refractive scintillation). Diffractive scintillation in this case is an unlikely explanation, as the calculated \( \Delta f_{\text{DISS}} \) is about \( 10^{-2} \) MHz, far lower than the observing frequency channel width \( \sim 1 \) MHz. For refractive scintillation, the calculated modulation index \( m = \sqrt{u^{5/3}} \sim 16\% \), and the timescale \( \sim 10^2 \) days. Walker (1998) estimated the effects of interstellar scintillation (arising from the Milky Way) on extragalactic sources using the TC93 model (Taylor & Cordes 1993), and we find a different modulation index of \( m \sim 42\% \) and timescale of \( \sim 20 \) days based on this model. These estimations are not too far from our observed results, and we therefore think the variability can be explained by propagation effects. However, we cannot rule out the possibility of contribution from intrinsic variation. As we mentioned before, the large duty cycle of the pulsar suggests that the pulsar could be an aligned rotator, which is known to show larger levels of nulling (Cordes & Shannon 2008). Aligned rotating pulsars also show mode-changing behavior (e.g., PSR J1107-5907, Hobbs et al. 2016).

The source shows very steep spectral indices (varying from \( \alpha \sim -4 \) to \( \alpha \sim -2 \)) even compared to the pulsar population (average \( \alpha \) of \(-1.4\); Bates et al. 2013, and \( \alpha < -2.5 \) for the fastest rotating millisecond pulsar, e.g., Frail et al. 2016a). The high-quality MeerKAT continuum data show an unusual radio spectral shape, with an upturn at 1.36 GHz from \( \alpha = -4.4 \) to \( \alpha = 2.34 \). The upturn at \( \sim 1 \) GHz is hard to explain as most pulsars can be described using a simple power-law spectrum. Some pulsar spectra are known to show a turnover at low frequencies due to synchrotron self-absorption or thermal free–free absorption; however, this is the opposite of what we see (i.e., transitioning from a flat to a steeper spectrum; see examples in Jankowski et al. 2018). We also note that some pulsars can have an upturn spectrum at a higher frequency, but this usually occurs at millimeter wavelengths (\( \gtrsim 10 \) GHz; Kramer et al. 1996). Therefore, we instead consider external effects. Tuntsov et al. (2017) found that scintillation can cause kinks, bumps, and wiggles in the broadband radio spectrum of a quasar. If PSR J0523–7125 is scintillating (as we suggested above), we might be able to see unusual spectra structures like these.

Further observations should be able to conclusively determine a timing solution, providing an age and spin-down luminosity of the source. We will discuss these properties in a later paper.

### 3.2. Identifying Pulsars in Continuum Surveys

The Magellanic Clouds have been targeted several times with pulsar surveys using the Parkes telescope (e.g., McCulloch et al. 1983; McConnell et al. 1991; Crawford et al. 2001; Manchester et al. 2006; Ridley et al. 2013), identifying 31 extragalactic pulsars. PSR J0523–7125 is brighter than all of these but was not identified in these surveys. Its wide pulse profile and steep spectrum could be responsible for its nondetection since they would reduce the signal-to-noise ratio. This is especially true in surveys with the Parkes Multibeam receiver at 1400 MHz, where many recent searches have been conducted. As described in Section 2.2, we cannot easily detect the pulsar using traditional methods from targeted Parkes data. However, we can clearly identify it in Parkes using only low-frequency data from the UWL receiver, consistent with the spectral properties determined from continuum images.

There are many continuum-based pulsar surveys using a steep spectral index as the selection metric (e.g., Damico et al. 1985; Camilo et al. 2000; Maan et al. 2018; Hyman et al. 2019), but only a limited number of new pulsars have been found (e.g., Bhakta et al. 2017; Frail et al. 2018). The steep-spectrum selection normally requires two observations at different frequencies to measure the spectral index. Moreover, there are also other types of sources that can have very steep spectra, such as high-redshift radio galaxies (O’Dea 1998). Circularly polarized sources without deep infrared/optical counterparts are highly likely to be pulsars. We also note that such sources may belong to an as yet unknown class or classes. For instance, the Galactic Center Radio Transient (GCRT; Hyman et al. 2002, 2005), two steep-spectrum, polarized sources near the Galactic bulge (C1748–2827 and C1709–3918; Hyman et al. 2021), and a polarized transient recently discovered in the VAST survey near the Galactic Centre (ASKAP J173608.2–321635; Wang et al. 2021) are circularly polarized sources without deep infrared/optical counterparts, whose natures are still unknown.

There are an increasing number of large-scale radio continuum surveys, ranging from low-frequency surveys such as the Low-Frequency Array (LOFAR) Two-meter Sky Survey (LoTSS; Shimwell et al. 2017), the Giant Metrewave Radio Telescope (GMRT) 150 MHz All-Sky Survey (TGSS; Intema et al. 2017), and the GaLactic and Extragalactic All-sky Murchison Widefield Array survey (GLEAM; Hurley-Walker et al. 2017), to gigahertz-frequency surveys such as the ASKAP Rapid Commissioning Survey (RACS; McConnell et al. 2020), the Evolutionary Map of the Universe survey with ASKAP (EMU; Norris et al. 2021), the Polarization Sky Survey of the Universe’s Magnetism (POSSUM; using the same sky coverage as EMU with full polarization; Gaensler et al. 2010), and the Karl G. Jansky Very Large Array (VLA) Sky Survey (VLASS; Lacy et al. 2020). With improved instruments in the Square Kilometre Array era, instantaneous large fields of view and great sensitivity will be even more common, leading to the detection of large numbers of radio sources across the sky. Some or most of the surveys include circular polarization measurements. Apart from radio wavelengths, there are significant improvements in multiwavelength surveys (with better sensitivity and large sky coverage), e.g., the VISTA Variables in the Via Lactea (VVV; Minniti et al. 2010), the VISTA Hemisphere Survey (VHS; McMahon et al. 2013), the Dark Energy Survey (DES; Abbott et al. 2018), and the Gaia mission (Gaia Collaboration et al. 2016). These deep multiwavelength surveys can greatly help with the classification of a circularly polarized object (e.g., to check if there is any stellar counterpart). As shown in Figure 7, which listed known circular polarized objects including various stellar
The near-infrared data were taken from various surveys with the Visible and Infrared Survey Telescope for Astronomy represents another polarized transient ASKAP J173608.2 et al. 2006.888 MHz dwarf BDR J1750 et al. 2016. objects, the radio to near-infrared flux ratio (plus the fractional circular polarization) can be used as a good diagnostic tool to select strong pulsar candidates.

One limitation of circular polarization selection is that not all pulsars have a high level of circular polarization (the median fraction \( \sim 10\% \); Han et al. 1998; Johnston & Kerr 2018). As estimated by Kaplan et al. (2019), even deep ASKAP continuum searches (e.g., EMU, with sensitivity \( \sim 50 \mu\text{Jy} \)) for circularly polarized sources are unlikely to find large numbers of new pulsars. However, it may be critical to identify extreme pulsars that are missed in traditional pulsar surveys. For example, the high electron density in the Galactic Centre direction (which has much higher stellar densities) causes strong scattering and makes pulsars relatively difficult to discover through traditional periodicity searches.

Extragalactic pulsar searches in continuum images are also possible. To date, the Magellanic Clouds are still the only place where extragalactic pulsars have been detected, even after several attempts at searching for pulsars in other galaxies including M31, M33, and nearby dwarf galaxies (e.g., McLaughlin & Cordes 2003; Bhat et al. 2011; Rubio-Herrera et al. 2013; Mikhailov & van Leeuwen 2016; van Leeuwen et al. 2020). We can estimate the number of detectable pulsars in M31 (our neighbor galaxy) through continuum surveys, by putting all known Galactic pulsars at the distance of M31 (785 \( \pm 25 \) kpc; McConnachie et al. 2005). We assume that the number of pulsars in M31 and our Galaxy is similar, though some work suggests there may be a smaller radio pulsar population in M31 (van Leeuwen et al. 2020). We consider using the Next Generation Very Large Array (ngVLA; Murphy et al. 2018), which is about 10 times more sensitive than the VLA and can achieve a resolution at \( \sim \) mas level. The ngVLA is located in the Northern hemisphere and therefore also ideal for observing M31 (and M33). The rms sensitivity of the ngVLA is expected to be \( \sim 0.1 \mu\text{Jy beam}^{-1} \) for a 10 hr observation at \( \sim 1–3 \) GHz. Assuming this sensitivity, more than \( \sim 50 \) simulated M31 pulsars could be detected at the \( >5\sigma \) level in a survey of M31, and \( \sim 5 \) could be detected in circular polarization. Although the estimated number is still very low, it is a promising possibility to detect the most luminous pulsars in M31 (e.g., those similar to PSR J0523–7125). With such an observation, we would also be able to measure their flux densities and spectral properties, providing valuable information for follow-up targeted periodicity searches.

4. Conclusion

We discovered a highly variable, circularly polarized object, VAST J052348.6–712552, in a variability analysis of two fields containing the Magellanic Clouds observed as part of the VAST-P1 survey. With \( \sim 20\% \) fractional circular polarization and no optical/infrared counterpart, VAST J052348.6–712552 was considered to be a strong pulsar candidate. Subsequent MeerKAT observations discovered a pulsar, PSR J0523–7125, associated with the continuum source, which was further confirmed with observations using the Parkes telescope.
pulsar has a period of 322.5 ms and a DM of 157.5 pc cm$^{-3}$, consistent with an LMC origin. The extragalactic distance makes PSR J0523–7125 among the most luminous known pulsars for steady-state emission; in particular, it is brighter than all known pulsars in the Magellanic Clouds at both 400 MHz and 1400 MHz. Despite its high luminosity, PSR J0523–7125 remained undetected in several LMC pulsar surveys, which we suspect is largely due to its wide pulse profile and/or steep spectral shape. The wide pulse profile also suggests that PSR J0523–7125 could be an aligned rotator. The strong variability is likely due to scintillation effects, though we cannot rule out the possibility of any intrinsic variation. We measured a large RM = +456 rad m$^{-2}$ for this pulsar, which is about twice the most extreme [RM] of LMC pulsars. A preliminary analysis shows that PSR J0523–7125 is likely embedded in an evolved HII region, which would be further strengthened if it turns out to be a young pulsar in ongoing timing analysis.

Our discovery highlights the possibility of identifying pulsars (especially nonstandard pulsars) from continuum images, particularly when circular polarization is combined with (largely archival) multiwavelength data. Improved next-generation radio telescopes and an increasing number of large-scale multiwavelength surveys will bring large amounts of data with great sensitivity and resolution, giving us an unprecedented opportunity to identify more pulsars (even for extragalactic pulsars farther than the Magellanic Clouds) via continuum images.

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