A PLANET IN AN 840 DAY ORBIT AROUND A KEPLER MAIN-SEQUENCE A STAR FOUND FROM PHASE MODULATION OF ITS PULSATIONS

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ABSTRACT

We have detected a 12\( M_{\text{Jup}} \) planet orbiting in or near the habitable zone of a main-sequence A star via the pulsational phase shifts induced by orbital motion. The planet has an orbital period of 840 ± 20 days and an eccentricity of 0.15. All known planets orbiting main-sequence A stars have been found via the transit method or by direct imaging. The absence of astrometric or radial velocity detections of planets around these hosts makes ours the first discovery using the orbital motion. It is also the first A star known to host a planet within 1\( r \) of the habitable zone. We find evidence for planets in a large fraction of the parameter space where we are able to detect them. This supports the idea that A stars harbor high-mass planets in wide orbits.

Key words: planets and satellites: detection – stars: oscillations – stars: variables: delta Scuti

1. INTRODUCTION

The Kepler Space Telescope is now the most successful planet-finding mission to date. Most of its thousands of planets (and planet candidates) have been found around cool stars (Mullally et al. 2015), while exoplanets orbiting A stars remain elusive. After the announcement of an additional 1284 Kepler exoplanets in 2016 May (Morton et al. 2016), still fewer than 20 A stars have confirmed planets.

Radial velocity observations of cool sub-giants, most notably via the “Retired A stars” project (Johnson et al. 2007), suggest that planet occurrence rates reach a maximum for stars of 1.9–2.1\( M_{\odot} \) (Reffert et al. 2015). This coincides remarkably with the masses of main-sequence A stars in the classical instability strip, where delta Scuti (\( \delta \) Sct) pulsators are common. The apparent planet deficit around these main-sequence stars can be explained as an observational selection effect, caused by problems in the application of the most successful planet-hunting methods to these types of stars. The transit method has difficulty because of the pulsational luminosity variations, which amount to several mmag, and because planets occupy wider orbits around A stars (Johnson et al. 2011; cf. Lloyd 2011), resulting in a lower transit probability. The radial velocity method, on the other hand, is particularly hindered by the nature of A-type spectra. A stars are typically fast rotators, with the mean of the equatorial rotational velocity distribution exceeding 100 km s\(^{-1}\) (Abt & Morrell 1995; Royer et al. 2007). Their high effective temperatures lead to fewer, shallower absorption lines, and these lines can be distorted by pulsation. Therefore the wavelengths of their spectral lines are not a precise standard of measure, and the state of the art is limited to 1–2 km s\(^{-1}\) precision (Becker et al. 2015).

Fortunately, the same pulsations that limit RV and transit surveys of A stars can themselves be used as precise clocks for the detection of orbital motion (Murphy et al. 2014). The pulsations of \( \delta \) Sct stars are particularly well-suited to this task (Compton et al. 2016).

1.1. Known Planets around A Stars

No planets have yet been discovered orbiting main-sequence A stars via the motion of their host stars, namely via the RV, astrometric, or pulsation-timing methods.

Discoveries have been made using two other methods: transits and direct imaging. The former category includes Kepler planet hosts (Kepler-13, 462, 516, 959, 1115, 1171, and 1517; Morton et al. 2016), plus WASP-33 (Collier Cameron et al. 2010), HAT-P-57 (Hartman et al. 2015), and KELT-17 (Zhou et al. 2016). The orbital periods of these systems are all under 30 days, except for Kepler-462 b, whose period is 85 days. These planets are strongly irradiated and would have high surface temperatures.

The imaged planets in the NASA Exoplanet Archive have very wide orbits. These include three objects (HD 100546b, HIP 78530b, and \( \kappa \) And b) with masses above the canonical 13–\( M_{\text{Jup}} \) boundary between planets and low-mass brown dwarfs, which orbit B-type stars. Exoplanets orbiting A stars include Fomalhaut b (Stapelfeldt et al. 2004; Chiang et al. 2009), \( \beta \) Pic b (Lagrange et al. 2010), HD 95086b (Rameau et al. 2013), the multi-planet system HR 8799c,e (Marois et al. 2008), and the planet in the triple system HD 131399 Ab (Wagner et al. 2016), all of which have orbits ranging from 9 to 177 au.

Finally, there are some planets orbiting evolved hot stars, such as those discovered from eclipse timing variations of the white-dwarf binary NN Ser (Brinkworth et al. 2006; Beuermann et al. 2010; Parsons et al. 2010), the companion to the sdB star V0391 Peg from pulsation-timing variations (Silvotti et al. 2007), and companions to other sdB stars KIC 10001893 (Silvotti et al. 2014) and the now-doubtful case of KIC 5807616 (Charpinet et al. 2011, Krzesinski 2015) from orbital brightness modulations.

In this paper we present the first planet discovered around a main-sequence A star by the motion of the host star, and the first to be in or near the habitable zone. We describe the observations and our method in Section 2. The planet’s habitability and the implications for planet occurrence posed by this discovery are discussed in Section 3.
2. OBSERVATIONS AND ANALYSIS

The host star, KIC 7917485, was observed for the full four years of the Kepler mission, in long-cadence mode (30 minute sampling). Our analysis used the multi-scale MAP data reduction of this data set (Smith et al. 2012; Stumpe et al. 2012). The time-series spans 1461 days, with a duty cycle of 91%.

This \( V = 13.2 \) mag star was observed spectroscopically with LAMOST (De Cat et al. 2015). The stellar atmospheric parameters were extracted by Frasca et al. (2016), showing KIC 7917485 to be a main-sequence A star, in agreement with the photometric characterization by Huber et al. (2014). Its position on a \( T_{\text{eff}}-\log g \) diagram is shown in Figure 1, and atmospheric parameters are given in Table 1. The star is located in the Sct instability strip and has a mass of approximately 1.63 \( M_{\odot} \).

We detected the planet by phase modulation (PM) of the stellar pulsations (Murphy et al. 2014). A Fourier transform of the light curve (Figure 2(a)) is dominated by two oscillation modes at frequencies \( f_1 = 15.3830026(1) \) and \( f_2 = 20.2628968(3) \) day\(^{-1} \), where the uncertainty on the final digit has been given in parentheses. These two modes were used in the PM analysis. Other significant peaks have amplitudes that are at least an order of magnitude lower. Their signal-to-noise ratios are too low to add usefully to the PM analysis, but their presence adds unwanted variance to the data. We therefore subtracted all peaks above 50 \( \mu \)mag from the data, except for \( f_1 \) and \( f_2 \), by fitting their frequencies to the light curve with a nonlinear least-squares algorithm. We also high-pass-filtered the light curve to remove any remaining instrumental signal and low-frequency oscillations, preserving all content at frequencies above 5 day\(^{-1} \). The Fourier transform after the additional processing is shown in Figure 2(c).

We divided the light curve into 10 day segments to look for shifts in the phases of \( f_1 \) and \( f_2 \). These phases were converted into delays in the light arrival time ("time delays") following the method of Murphy et al. (2014). Time delays of \( f_1 \) and \( f_2 \) show identical periodic variation (Figure 3), which we attribute to a sub-stellar companion. Values for the orbital parameters were initially obtained using formulae from Murphy & Shibahashi (2015), and then refined with an MCMC algorithm. The MCMC analysis used a Metropolis–Hastings algorithm (Metropolis et al. 1953; Hastings 1970) with symmetric proposal distributions based on Gaussian-distributed random numbers. Trial runs were made to determine appropriate step
sized in each of the five orbital parameters fitted: the orbital period, $P_{\text{orb}}$, projected light travel time across the orbit, $a_1 \sin i / c$, eccentricity, $e$, the phase of a pair of passage calculated relative to the first time-delay observation, $\phi_p$, and the angle of the ascending node, $\varpi$.

4.2. Implications for Planet Occurrence Rates

KIC 7917485b is the least massive companion that we have found in *Kepler* data with the PM method (S. J. Murphy et al. 2016, in preparation). We have also found two other stars with time delay variations consistent with planetary companions having periods longer than the 4-year data set ($KIC\ 9700322$, $KIC\ 8453431$), but the finite duration of *Kepler* time-series does not allow the orbits to be fully parametrized. These detections allow us to comment on the planet occurrence around A stars.

The detectability of low-mass companions is very sensitive to the noise in the Fourier transform of the time delays, which is determined by the pulsation properties (Murphy et al. 2016). We quantified this noise level for 2040 pulsating single A stars and found that only 5 of them had lower noise levels than KIC 7917485. Against the same sample, KIC 9700322 and KIC 8453431 ranked as the 9th lowest and 2nd lowest, respectively. In other words, we have been able to detect a planetary-mass companion in one of the nine stars with the lowest noise levels, and two others show variations that are consistent with planetary-mass companions. This fact is in strong support of existing observations that intermediate-mass stars ("retired A stars") tend to host high-mass planets in wide orbits (Johnson et al. 2011).

Table 2 gives the best-fitting orbital parameters, which are compared with the observations in Figure 4.

Table 2
Orbital Parameters for the KIC 7917485 System

| Parameter      | Units          | Values   |
|----------------|----------------|----------|
| $P_{\text{orb}}$ | days          | 840$^{+22}_{-20}$ |
| $e$            |               | 0.15$^{+0.15}_{-0.10}$ |
| $\phi_p$      | [0–1]         | 0.89$^{+0.010}_{-0.012}$ |
| $f(m_1, m_2, \sin i)$ | $M_\odot$ | 5.3$^{+0.07}_{-1.5}$ $\times 10^{-7}$ |
| Star           |               |          |
| $M_1$         | $M_\odot$     | 1.63$^{+0.15}_{-0.12}$ |
| $a_1 \sin i / c$ | s             | 7.1$^{+0.5}_{-0.4}$ |
| $\varpi$      | rad           | 3.0$^{+0.06}_{-0.07}$ |
| $K_1 \sin i$  | m s$^{-1}$    | 186$^{+17}_{-13}$ |
| Planet         |               |          |
| $M_2 \sin i$  | $M_\odot$     | 0.0113$^{+0.0008}_{-0.0006}$ |
| $M_{\text{upp}}$ |               | 11.8$^{+0.8}_{-0.6}$ |
| $a_2 \sin i / c$ | s             | 1017$^{+136}_{-140}$ |
| $a_2 \sin i$  | au            | 2.03$^{+0.22}_{-0.22}$ |
| $\varpi$      | rad           | 6.1$^{+0.7}_{-0.7}$ |

Note. The value of $\phi_p$ is calculated with respect to the first time-delay measurement at $t_0$ (BJD) = 2,454,958.39166. $M_1$ is inferred non-dynamically, from the spectroscopic parameters, which allows $M_2 \sin i$ and $a_2 \sin i$ to be calculated. $K_1 \sin i$, the projected radial velocity semi-amplitude, is calculated from the orbital parameters and provided for reference; it is not used to derive the orbital solution.

Figure 4. Comparison of the weighted average time delays with the best-fitting orbital parameters from the MCMC analysis (given in Table 2).

The mean projected separation between the components, $a \sin i = a_1 \sin i + a_2 \sin i$, is 2.05 au. At this separation, the luminous A star irradiates the companion to a surface flux ratio, $S/S_0$, of 2.36 times the flux at Earth. Statistical correction for random inclination gives $S/S_0 = 1.76$ as the most probable value.

The luminosity of the A star is not well-constrained, and is a strong function of main-sequence age. At the $1\sigma$ lower-luminosity limit (see Table 1), those values of $S/S_0$ reduce to 1.72 and 1.27, respectively. Thus the position of KIC 7917485b is consistent with the habitable zone at the $1\sigma$ level. The companion would have been closer to the center of the habitable zone earlier in the star’s lifetime. The luminosity will be refined substantially when the *Gaia* mission provides a distance measurement, and the question of habitability can be reassessed.

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4. CONCLUSIONS

We have found a planetary companion of $M \sin i = 11.8$ $M_{\text{upp}}$ near the inner edge of the habitable zone of KIC 7917485. This is the first planet orbiting a main-sequence A star to be discovered via the motion of its host star. Other planets orbiting A stars have been discovered via the transit method and have short periods, or been discovered by direct imaging and have very long periods. Our finding is particularly significant because no other method is presently capable of detecting...
non-transiting planets around these stars with periods of a few years, i.e., near their habitable zones. KIC 7917485 has particularly low noise levels in its light arrival time delays. We analyzed other stars with similarly low noise and for two of them we also found evidence for planetary-mass companions with periods similar to the 4-year time span of Kepler data. This fact strongly supports the idea that intermediate-mass stars tend to host high-mass planets in wide orbits.

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