First Experimental Confirmation of the CH$_3$O + H$_2$CO $\rightarrow$ CH$_3$OH + HCO Reaction: Expanding the CH$_3$OH Formation Mechanism in Interstellar Ices

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Abstract

The successive addition of H atoms to CO in the solid phase has been hitherto regarded as the primary route to form methanol in dark molecular clouds. However, recent Monte Carlo simulations of interstellar ices alternatively suggested the radical-molecule H-atom abstraction reaction CH$_3$O + H$_2$CO $\rightarrow$ CH$_3$OH + HCO, in addition to CH$_3$O + H $\rightarrow$ CH$_3$OH, as a very promising and possibly dominating (70%–90%) final step to form CH$_3$OH in those environments. Here, we compare the contributions of these two steps leading to methanol by experimentally investigating hydrogenation reactions on H$_2$CO and D$_2$CO ices, which ensures comparable starting points between the two scenarios. The experiments are performed under ultrahigh vacuum conditions and astronomically relevant temperatures, with H:$H_2$CO (or D$_2$CO) flux ratios of 10:1 and 30:1. The radical-molecule route in the partially deuterated scenario, CHD$_2$O + D$_2$CO $\rightarrow$ CHD$_2$OD + DCO, is significantly hampered by the isotope effect in the D-abstraction process, and can thus be used as an artifice to probe the efficiency of this step. We observe a significantly smaller yield of D$_2$CO + H products in comparison to H$_2$CO + H, implying that the CH$_3$O-induced abstraction route must play an important role in the formation of methanol in interstellar ices. Reflection-absorption infrared spectroscopy and temperature-programmed desorption-quadrupole mass spectrometry analyses are used to quantify the species in the ice. Both analytical techniques indicate constant contributions of $\sim$80% for the abstraction route in the 10–16 K interval, which agrees well with the Monte Carlo calculations. Additional H$_2$CO + D experiments confirm these conclusions.

Unified Astronomy Thesaurus concepts: Astrochemistry (75); Laboratory astrophysics (2004); Dark interstellar clouds (352)

1. Introduction

Methanol (CH$_3$OH) is abundantly detected in interstellar environments and is one of the main components of interstellar ices. Its formation is tightly connected to the evolution of molecular clouds. In these dense and cold regions, CO molecules present in the gas-phase freeze-out and form an apolar coating on top of icy grains, largely comprising of H$_2$O and CO$_2$ ice. Methanol is then efficiently formed through the hydrogenation of CO in the solid phase (Tielens & Hagen 1982; Charnley et al. 1992; Hiraoka et al. 1994; Watanabe & Kouchi 2002; Fuchs et al. 2009). The presence of methanol in CO-rich ice is supported by a series of comparisons between laboratory and observational evidence (e.g., Bottinelli et al. 2010; Cuppen et al. 2011; Penteado et al. 2015). Furthermore, as opposed to CO hydrogenation, neither gas-phase nor alternative solid-phase routes can explain the observed CH$_3$OH abundances (Geppert et al. 2005; Garrod et al. 2006; Watanabe et al. 2007). As a secondary mechanism, it has also been shown to form in H$_2$O-rich ices prior to the heavy CO freeze-out—thus in an earlier evolutionary stage of the cloud (Bergner et al. 2017; Qasim et al. 2018; Molpeceres et al. 2021; Potapov et al. 2021).

So far, the main proposed hydrogenation route to form methanol is through successive addition reactions of H atoms to CO ice, as follows:

\begin{align*}
\text{CO} + \text{H} & \rightarrow \text{HCO} \quad (1) \\
\text{HCO} + \text{H} & \rightarrow \text{H}_2\text{CO} \quad (2) \\
\text{H}_2\text{CO} + \text{H} & \rightarrow \text{CH}_3\text{O} \quad (3) \\
\text{CH}_3\text{O} + \text{H} & \rightarrow \text{CH}_3\text{OH}. \quad (4)
\end{align*}

Here, abstraction reactions induced by H atoms also take place, converting CH$_3$OH into CO through H$_2$CO (Hidaka et al. 2009; Chuang et al. 2016). However, recent theoretical works that combine a full reaction network for the hydrogenation of CO and microscopic kinetic Monte Carlo simulations of interstellar ices suggest a dominating alternative to Reaction (4), which is the radical-molecule route (Álvarez-Barcia et al. 2018; Simons et al. 2020):

\begin{align*}
\text{CH}_3\text{O} + \text{H}_2\text{CO} & \rightarrow \text{CH}_3\text{OH} + \text{HCO}, \quad (5)
\end{align*}

with a contribution of 70%–90% with respect to the remainder of the routes to form methanol. Reaction (5) can take place upon only one hydrogen-atom addition to H$_2$CO, as long as another H$_2$CO molecule lies in the vicinity of the formed CH$_3$O. Comparatively, Reaction (4) warrants that an additional H atom diffuses through the ice and reacts with CH$_3$O. The overall contribution from each route will therefore be dictated by the availability of H, CH$_3$O, and H$_2$CO in the ice. Although the theoretical results represent a significant paradigm change in the final step of the CO hydrogenation formation pathway of methanol in interstellar ices, they have
not yet been systematically investigated in the laboratory. Experimental verification is however possible by exploring the kinetic isotope effect of the reactions involving deuterated species (i.e., D₂CO and D), as will be discussed below.

At the typical low temperatures of molecular clouds, quantum-mechanical tunneling governs the activation of chemical reactions. Under these conditions, D-abstraction reactions are significantly hampered in comparison to their H-abstraction counterparts (e.g., Nagaoka et al. 2005; Goumans et al. 2007; Hidaka et al. 2011), on account of the so-called “kinetic isotope effect.” Reaction (5) is a radical-molecule route that involves a H-abstraction reaction from H₂CO with an activation barrier of ΔE = 2670 K (Álvarez-Barcia et al. 2018). Thus, it most likely proceeds through quantum tunneling at such low temperatures. As a consequence, the rate constant of the analog reaction with the deuterated formaldehyde (i.e., D-atom abstraction from D₂CO) is expected to be severely hindered. In the present work, we take advantage of this isotope effect to quantify the respective contributions from mechanisms (4) and (5) in forming methanol under molecular-cloud conditions. We then confirm the dominance of the radical-molecule route through complementary H₂CO + D experiments.

2. Experimental

All experiments are performed using the setup SURFRESID3, which has been described in detail elsewhere (Ioppolo et al. 2013; Qusim et al. 2020). It consists of an ultrahigh vacuum chamber with a base pressure of ~5 × 10⁻¹⁰ mbar. At its center, a gold-plated copper substrate is mounted on the tip of a closed-cycle He cryostat that allows the temperature of the substrate to vary between 8 and 450 K. The substrate is heated at a ramping rate of 5 K minutes⁻¹, and similar rate constants are not performed due to the low efficiency of the D-atom abstraction from D₂CO (Tschersich 2000). Additionally, analogous codeposition experiments are performed at temperatures of 10, 12, 14, and 16 K. Furthermore, complementary set of H₂CO + D and D₂CO + D codepositions are not performed due to the low efficiency of the D-atom addition reactions resulting from the kinetic isotope effect caused by the slower rate of quantum tunneling of D atoms over H atoms (Hidaka et al. 2009; Goumans 2011).

During the experiments, the ices are monitored in situ by means of Fourier-transform reflection-absorption infrared spectroscopy (RAIRS) in the range of 700–4000 cm⁻¹, with a resolution of 1 cm⁻¹. After the codeposition experiments, the sample is heated at a ramping rate of 5 K minutes⁻¹. The desorbed species are monitored by a quadrupole mass spectrometer during temperature-programmed desorption experiments (TPD-QMS). To derive the column densities of the species in the ice, we employ the modified Beer–Lambert law to convert absorbance area to absolute abundance. We use band-strength values of A(H₂CO) = 1727 × 3.6 × 10⁻¹⁷ cm molecule⁻¹ and A(CH₃OH) = 1030 × 3.1 × 10⁻¹⁷ cm molecule⁻¹, which were calibrated in the same setup by laser-interference measurements (Chuang et al. 2018). For CHD₂OH (ν = 1037 cm⁻¹), the band strength is estimated by multiplying A(CH₃OH) by a calibration factor of 0.43 (Nagaoka et al. 2007). Finally, the band strength of D₂CO is obtained by comparing the infrared (IR) absorbance areas of the same amount of pure H₂CO and D₂CO ices. We experimentally derive A(D₂CO) = 1679 × 2.8 × 10⁻¹⁷ cm molecule⁻¹. More details on this procedure are given in Appendix A. For a specific flux configuration, the relative uncertainties of the molecule and H-atom fluxes are, respectively, <9% and <4%. The error bar estimation due to the instrumental uncertainties of the RAIRS and TPD-QMS measurements is derived from the integrated noise signals of blank experiments in the same bandwidth. Table 1 in Appendix B lists the experiments performed in this study and the relative abundance of CHD₂OH/CH₃OH derived from the RAIRS and QMS-TPD analyses.

3. Results and Discussion

In Figure 1, we present a schematic view of the reaction routes probed in our experiments. To focus on Reactions (4) and (5), we skip steps (1) and (2) and start by directly depositing formaldehyde. This way, comparable starting points between the two scenarios are obtained.

By using different isotope-labeled reactants (H₂CO and D₂CO), the importance of Reactions (4) and (4') can be directly assessed. The yields of methanol from H₂CO through Reaction (4) and from D₂CO through (4')—see Figure 1—should be identical, since both radical–radical routes are barrierless. Conversely, the formation of the isotope-labeled CHD₂OD through Reaction (5') is expected to be strongly affected by the isotope effect and to become negligible, whereas (5) still takes place. Moreover, the hydrogen-addition steps H₂CO + H and D₂CO + H have similar activation barriers of 2318 and 2253 K, respectively, and similar rate constants (Hidaka et al. 2009; Goumans 2011). Thus, they yield comparable amounts of CH₃O and CHD₂O. The contributions of the two competing routes are hence obtained by comparing the relative abundances between the total CH₃OH yield of the H₂CO ice (Reactions (4) and (5)) and that of CHD₂OH in the D₂CO ice (Reaction (4')).

The final infrared spectra of both H₂CO + H and D₂CO + H codepositions after 360 minutes at 12 K are shown in Figure 2. IR features of H₂CO are observed at 1167 (CH₂ wag., ν = 6), 1249 (CH₂ rock., ν = 5), 1500 (CH₂ scis., ν = 3), and 1727 cm⁻¹ (CO str., ν = 2). In the analog deuterated experiments, D₂CO peaks are detected at 988 (CD₂ wag.), 1101 (CD₂ scis.), and 1679 cm⁻¹ (CO str.). In both cases the hydrogenated products CH₃OH and CHD₂OH are detected by their CO stretching (ν = 8) features at 1030 and 1037 cm⁻¹, respectively (Shimanouchi 1972; Hidaka et al. 2009), which indicates that the H-addition routes in Figure 1 proceed at a considerable rate.

Even in the experimental configuration with the highest product yield, no CHD₂OD peaks (OD str. at ν = 2420 cm⁻¹, Nagaoka et al. 2007) could be observed in the D₂CO + H spectrum. The detection of CHD₂OH and nondetection of CHD₂OD are in line with the expected low efficiency of the
radical-molecule route for deuterated species and are consistent with the results by Hidaka et al. (2009) on the exposure of H atoms to D2CO ice. That, in addition to the fact that the CHD2OH peak is much smaller than the CH3OH counterpart in the H2CO + H experiment at equivalent conditions, strongly suggests that Reaction (4) is not the dominant route to form methanol under the investigated experimental conditions. Instead, Reaction (5) must play a more prominent role in the formation of methanol. This observation remains valid after converting the integrated peak areas from the RAIIR spectra to the species’ abundances, as shown in Table 1. From the RAIIR analyses, we do not detect any bands related to the radicals HCO or CH2OH—nor their deuterated isotopes—reinforcing the findings from previous studies that their high reactivity results in low abundances (Watanabe & Kouchi 2002; Fuchs et al. 2009). The radical CH3O is also not observed in the H2CO + H experiments. It should be noted, though, that the most intense IR features of its isotopologue CHD2O overlap with those of the other reactants and products in the spectra (see, e.g., Haupa et al. 2017), and therefore we cannot rule out its presence in the D2CO + H experiments. Moreover, we observe a clear CO peak (ν = 2138 cm−1) in the H2CO + H experiment, as well as HDCO (ν = 1694 cm−1) and H2CO (ν = 1727 cm−1) peaks in the D2CO + H counterpart, in agreement with the results by, e.g., Hidaka et al. (2009) and Chuang et al. (2016) that H-atom-induced abstraction reactions take place for the studied conditions (see, e.g., Figure 5 in Hidaka et al. 2009).

The TPD-QMS results of the experiments at 12 K are presented in Figure 3 as a typical example of the obtained spectra. Since trying to account for the mass-fragmentation patterns induced by electron impact may complicate the interpretation of the data, only the ion signals of the products’ molecular masses are selected to quantify their formation yield. For example, the CH3OH formation is quantified by the mass signal m/z = 32, while the deuterated species CHD2OH and CHD2OD are depicted by m/z = 34 and 35, respectively. All methanol (−d0−d3) mass signals peak at ∼140 K. The small m/z = 32 signal that peaks at 160 K could be due to either CH3OH trapped in background water or fragments originating from complex organic molecules (He et al. 2022; Fedoseev et al. 2015; Chuang et al. 2016). The flux of background water deposited throughout the experiment is estimated to be 5.2 × 1010 molecules cm−2 s−1, being a minor component of the ice that does not affect the final results. The TPD spectra show that a much smaller yield of CHD2OH is observed in comparison to CH3OH, and no CHD2OD is detected above the instrumental detection limit, thereby confirming the RAIIR analysis.

Figure 4 presents the relative yields of CH3OH and CHD2OH derived from the RAIIR and QMS-TPD analyses at 10, 12, 14, and 16 K. The methanol yield of the hydrogenation reactions varies with the surface temperature, showing the effective result of the competition between the increase in the diffusion of H atoms and the decrease in their residence time on the ice as a function of temperature from 10 to 16 K. The highest abundances of methanol in both experiments H2CO + H and D2CO + H are found at 12 K, and drop with increasing temperatures. A set of experiments performed at 20 K results in CHD2OH yields that are not measurable, and thus are not reported here. Therefore, the abundances are normalized to that of CH3OH at 12 K. Similar observations have been reported in a previous study on the hydrogenation of H2CO at temperatures up to 25 K (Chuang et al. 2016).

For the entire set of experiments, the RAIIR data show a considerably lower abundance of CHD2OH from the D2CO + H experiments compared to that of CH3OH from H2CO + H. In fact, the abundance ratios of CHD2OH/CH3OH between analogous experiments yield a constant value of ~0.16 for the temperature range between 10 and 16 K (see Table 1). This is confirmed by the QMS-TPD data, which agree well within their uncertainties with the results from the RAIIR analysis (CHD2OH/CH3OH ~0.17). The contributions of the reactions probed here are therefore independent of these temperatures (10–16 K), which are of relevance to molecular clouds. This observation is in accordance with the predictions of the models by Simons et al. (2020). The similarity between

\[
\text{H}_2\text{CO} + \text{H} \rightarrow \text{CH}_3\text{O} + \text{H} \quad (4)
\]

\[
\text{D}_2\text{CO} + \text{H} \rightarrow \text{CHD}_2\text{O} + \text{H} \quad (4')
\]

\[
\text{CH}_3\text{OH} + \text{H} \rightarrow \text{CH}_3\text{O} + \text{H} \quad (5)
\]

\[
\text{CH}_3\text{OH} + \text{H} \rightarrow \text{CH}_3\text{O} + \text{H} \quad (5')
\]
both the RAIRS and TPD ratios also excludes the possibility of the former being influenced by the heating of the substrate.

Under the assumption that Reactions (3) and (3') result in similar amounts of CH$_3$O and CHD$_2$O, respectively, and since both Reactions (4) and (4') are barrierless, the methanol yields from the abstraction route should be similar in both H$_2$CO + H and D$_2$CO + H experiments;

\[(4) \sim (4'), \quad (6)\]

where the parentheses denote the methanol yield of the reactions. The discrepancy between the yields of CH$_3$OH and CHD$_2$OH under our experimental conditions can therefore be used to estimate the contribution C4 of Reaction (4) with respect to the total yield of methanol:

\[
C_4 = \frac{(4)}{(4) + (5)} \sim \frac{(4')}{(4) + (5)} \\
\sim \frac{(\text{CHD}_2\text{OH})}{(\text{CH}_3\text{OH})} \sim [0.16 - 0.17]. \quad (7)
\]

The estimation described here assumes that the formation of CHD$_2$OH and CH$_3$OH takes place predominantly through the reactions presented in Figure 1. Other routes to form methanol, such as CH$_3$OH + H, have very minor contributions and are hence disregarded based on the inefficient formation of CH$_3$OH from H$_2$CO hydrogenation (Song & Kästner 2017 and, e.g., Table 6 in Simons et al. 2020). Accordingly, the contribution C5 of Reaction (5) corresponds to:

\[C_5 = 1 - C_4 \sim [0.83 - 0.84]. \quad (8)\]

The experimentally derived value of C5 agrees well with the contribution of \(~90\%\) reported by Simons et al. (2020) from kinetic Monte Carlo simulations of the hydrogenation of H$_2$CO ices.

The H/H$_2$CO flux ratios employed here were chosen based on the conditions of the models by Simons et al. (2020) (H: H$_2$CO = 20:1). A higher H flux could potentially favor the atom-addition route and thus lead to a lower product ratio. To verify whether the contributions C4 and C5 depend on the hydrogen-to-molecule ratio, a set of experiments using H/H$_2$CO(D$_2$CO) = 30 was performed at 10 K (see Table 1). The resulting CHD$_2$OH/CH$_3$OH ratios derived from the RAIRS and TPD-QMS data are 0.18 ± 0.03 for C4 and 0.82 ± 0.03 for C5. These values are negligibly different considering the detection error. Thus, a higher hydrogen flux affects only slightly the contributions derived above, and Reaction (5) still governs the formation of CH$_3$OH. This implies that both flux ratios employed here represent an overabundance of H atoms, and therefore the product yield is limited by either H$_2$CO or D$_2$CO.

In our experiments, the dominance of the radical-molecule route is likely due to the higher availability of H$_2$CO in the environment, as opposed to H atoms, to react with CH$_3$O. Both CH$_3$O and H are very reactive species; when a CH$_3$O radical is formed in the vicinity of a H$_2$CO molecule, it readily reacts with H$_2$CO to form CH$_3$OH + HCO. On the other hand, the accreted H atoms on the surfaces are required to diffuse in order to react with the CH$_3$O radicals. Therefore, they will mostly react with the broadly available H$_2$CO, and less likely with CH$_3$O, which is a minor ice component. Moreover, the diffusion of hydrogen atoms also competes with its reaction with CH$_3$O, whereas both CH$_3$O and H$_2$CO do not diffuse and therefore have more time available to react. As a result, the higher likelihood of the two reactants in Reaction (5) to meet, overpowers Reaction (4), despite the latter being barrierless.

Although the ice composition explored here is not fully representative of interstellar ices, the foregoing reaction routes are nonetheless relevant to their chemistry. In more realistic ices, there is a lower availability of formaldehyde, and the formed CH$_3$O radicals can also react with other species. Nevertheless, the aforementioned reasoning still applies, since the H atoms will mostly react with CO and HCO, additionally to H$_2$CO. Indeed, route (5) was observed to dominate the formation of methanol even in simulations starting from the
hydrogenation of CO ice (Simons et al. 2020). Moreover, the radical-molecule route forms HCO as a byproduct, which will then be hydrogenated again to H2CO, replenishing and thus favoring the chemical network. Additional formaldehyde formed in the gas phase in the interstellar medium can also subsequently accrete onto ice grains and take part in the radical-molecule route.

It is noteworthy that the hydrogenation of pure D2CO ice is not perfectly equivalent to that of pure H2CO ice. The rate constant of Reaction (3)’ is in fact slightly higher than that of Reaction (3) (Hidaka et al. 2009; Goumans 2011). Moreover, Reaction (5)’ is inhibited by the isotope effect. Therefore, there should be an increase in the availability of CHD2O in the ice during the deuterated experiments, which in turn favors the formation of CHD2OH through route (4’). The surplus of CHD2O is either consumed through H-induced abstraction reactions or preserved in the ice. Furthermore, the isotope effect also hampers the D-abstraction reaction D2CO + H → DCO + HD, which is why the formation of CO is much less efficient in these scenarios (see Figure 2).

This results in a higher availability of H atoms in the surface of the ice in the deuterated experiments, compared to the nondeuterated counterparts where the abstraction reaction H2CO + H → HCO + H2 is faster. As a consequence, the formation of CHD2OH is also favored by the surplus of hydrogen atoms. Therefore, the contribution from Reaction (4’) to forming methanol under the present experimental conditions should be regarded as an upper limit to that of Reaction (4), which puts a lower limit to Reaction (5) as being at least roughly four times more efficient than Reaction (4).

Aiming to directly confirm the occurrence of the radical-molecule route forming methanol, an experiment of H2CO + D and a control experiment of D-atom beam exposure to the bare substrate are performed at 10 K for 360 minutes. The codeposition experiment is followed by a TPD with a ramping rate of 5 K/min. These experiments are designed to probe the following competing reactions:

\[
\text{H}_2\text{CO} + \text{D} \rightarrow \text{CH}_2\text{DO} + \text{D} \rightarrow \text{CH}_2\text{DOD} \tag{9}
\]

\[
\text{H}_2\text{CO} + \text{D} \rightarrow \text{CH}_2\text{DO} + \text{H}_2\text{CO} \rightarrow \text{CH}_2\text{DOH} + \text{HCO} \tag{10}
\]

The difference of the spectra acquired after the H2CO + D experiment and the D deposition is shown in panel (a) of Figure 5. In agreement with Hidaka et al. (2009), we observe spectral features of HDCO and D2CO as a result of D-substitution (abstraction + addition) reactions on H2CO. Moreover, a signal at \(\nu = 1034\ cm^{-1}\) is in line with the CO stretching band of CH2DOH, which would indicate that the radical-molecule route (i.e., Reaction (10)) proceeds. However, this feature is (at least partially) blended with the CH2 rocking mode of HDCO (\(\nu = 1029\ cm^{-1}\)), to an extent where deconvolution is not possible. Nonetheless, given the lower desorption temperature of HDCO (\(~100\ K\)) with respect to that of CH2DOH (\(~140\ K\)), it is possible to isolate the contribution from CH2DOH by analyzing the infrared spectra obtained during the TPD experiment. The spectra acquired at 10 K and in the temperature intervals of 80–120 K and 120–160 K are shown in panel (b) of Figure 5.

In the IR spectrum acquired in the temperature range of 80–120 K, the peak at \(\nu \sim 1029\ cm^{-1}\) contains the contributions from both HDCO and CH2DOH. However, at the interval of 120–160 K, the peak slightly shifts to \(\nu \sim 1034\ cm^{-1}\), and only accounts for the CH2DOH left on the substrate. Based on the obtained area of this peak and its reported band strength (Nagaoka et al. 2007), we derive a CH2DOD column density of \(~(3.0 \pm 0.3) \times 10^{14}\ \text{molecules cm}^{-2}\). This yield is ~4 times smaller than that of CH2OH from the H2CO + H experiment at 10 K (H/H2CO = 10), which is consistent with a lower rate for the D-addition reaction to H2CO over the H-addition counterpart.

During TPD, the mass signal \(m/z = 33\) is detected at ~150 K. However, whereas the spectroscopic signals are quite unambiguous, this does not apply to the mass spectrometric data. Since the mass signals of CH2DOH (\(m/z = 33, 32, 31, \ldots\)) overlap with those of CH2DOD (\(m/z = 34, 33, 32, \ldots\)), and since their fragmentation patterns are unknown, the interpretation of the mass spectrometry data is impeded.

The product of the D-addition route, CH2DOD (Reaction (9)), was not observed in its OD stretching mode (\(\nu \sim 2430\ cm^{-1}\)) within the detection limit of the instrument, as shown by the dashed line in panel (a) of Figure 5. Given that the band strength of this feature is expected to be around 1.5 times that of the CO stretching mode of CH2DOH (Nagaoka et al. 2007), the abundance of CH2DOD must be appreciably smaller than that of CH2DOH. This result provides additional compelling evidence that the radical-molecule route dominates the formation of methanol under the present conditions.

### 4. Conclusions

In the present work, we experimentally examine the contribution from the conventional route to the formation of methanol (Reaction (4)) with Reaction (5), and provide the respective impact of the two proposed last steps of relevance to the chemistry of molecular clouds. This was done by focusing on the specific reaction step in question, as more realistic experiments with CO ices would be much more difficult to interpret. Our main conclusions are the following:

1. We derive a contribution of \(\geq 83\%\) for the radical-molecule route (5), which is independent of the H/H2CO ratios (i.e., 10 and 30) and the temperatures (i.e., 10–16 K) experimentally investigated in this work. This temperature range is typical of dense clouds, where H atoms are expected to be overabundant in comparison with H2CO. Our experiments confirm the results in Simons et al. (2020), who performed simulations for conditions as used in the laboratory work presented here. Their model, in turn, extends to the interstellar medium, and concludes that this route dominates the formation of interstellar methanol. Our results corroborate with this finding. Comparatively, Reaction (4) should account for \(\lesssim 17\%\) of the methanol formation under the present conditions.

2. Additional codeposition experiments of H2CO and D yield measurable amounts of CH2DOH, whereas CH2DOD is not detectable within the IR experimental limits. This further confirms the dominance of the radical-molecule route to form methanol under our experimental conditions.

3. The conclusion drawn here does not affect the consensus that methanol is mainly formed through the hydrogenation of CO in the solid phase, albeit by means of a different mechanism than originally thought. Nonetheless, this finding represents an important change in the
perspective of methanol formation in space, and could affect astrochemical models that involve this key species.

4. The expected deuterium fractionation of methanol will likely be affected, with a possibly smaller abundance of deuterated species due to the kinetic isotope effect.

5. Furthermore, the HCO radical is an important astrochemical intermediate since it can react to further form H$_2$CO or recombine with other radicals (e.g., CH$_3$O and CH$_2$OH) to produce complex organic molecules (COMs). The alternative route probed here forms HCO radicals as a byproduct, which enriches an icy surface with this species and therefore affects the COM distribution in interstellar ices, likely enhancing the predicted formation rate of, e.g., glycolaldehyde and ethylene glycol in previous astrochemical models that did not consider Reaction (5) as a dominant channel. Models of the radical–radical reactions driven by the warming of the ice at later stages of star formation will also be affected by the additional HCO that eventually remains preserved in the ice and that becomes mobile and reacts.

6. The inclusion of deuterium atoms to microscopic kinetic Monte Carlo simulations would be a logical next step, as it could provide additional insight on the chemical network involving methanol. This, however, is beyond the scope of the present work.

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Appendix A

D$_2$CO Band Strength Estimation

The column densities ($N_X$) of the species in the ice are calculated by means of the modified Beer–Lambert equation:

$$N_X = \ln 10 \int \frac{Abs(\nu) d\nu}{A_X},$$

(A1)

in which $\int Abs(\nu) d\nu$ is the integrated absorbance of a given infrared band and $A_X$ is its corresponding band strength. To estimate the band strength of D$_2$CO, the same deposition conditions (i.e., exposure time and dose) are applied to both pure H$_2$CO and D$_2$CO, ensuring identical abundances of both types of ices. This is achieved by using the same leak valve while keeping the same flow-rate setting. Previously, Nagaoka et al. (2007) used this method to determine the relative integrated band strengths of CH$_3$OH and isotopes. We plot the integrated areas of the D$_2$CO $\nu_2 = 1679$ cm$^{-1}$ band versus those of H$_2$CO $\nu_2 = 1727$ cm$^{-1}$ for different accumulated abundances performed at three different flux configurations, as shown in Figure 6.

After performing a linear fit to the points in the plot, we derive an average slope of $a = 0.78 \pm 0.03$. This slope corresponds to the conversion factor between $A$(H$_2$CO) and $A$(D$_2$CO):

$$A(D_2CO)_{\nu_2} = a \times A(H_2CO)_{\nu_2}.$$  

(A2)

Given that $A$(H$_2$CO) for the $\nu = 1727$ cm$^{-1}$ peak has been previously measured to be $3.6 \times 10^{-17}$ cm molecule$^{-1}$ (Chuang et al. 2018), we estimate the D$_2$CO $\nu = 1679$ cm$^{-1}$ band strength to be $A$(D$_2$CO)$\sim 2.8 \times 10^{-17}$ cm molecule$^{-1}$.
Figure 6. Plot of the integrated area of the D$_2$CO ($\nu_2 = 1679$ cm$^{-1}$) infrared absorption band vs. that of H$_2$CO ($\nu_2 = 1727$ cm$^{-1}$) for three different molecule fluxes, $f$: $8 \times 10^{11}$, $1.6 \times 10^{12}$, and $3.2 \times 10^{12}$ cm$^{-2}$ s$^{-1}$.

Appendix B
Experiment List

In Table 1 below the experiments performed in this work are summarized.

Table 1
Overview of the Experiments Performed in This Work, as well as the Resulting CHD$_2$OH/CH$_3$OH Abundance Ratios between Each Set of Experiments Derived from the RAIRS and TPD-QMS Analyses

| Experiment | $T_{\text{sample}}$ (K) | Molecule Flux (cm$^{-2}$ s$^{-1}$) | H (D) Flux (cm$^{-2}$ s$^{-1}$) | Time (minutes) | CHD$_2$OH/CH$_3$OH (RAIRS, $\nu = 8/\nu = 8$) | CHD$_2$OH/CH$_3$OH (TPD) |
|------------|-----------------|-------------------------------|-----------------|-------------|-----------------------------|-----------------------------|
| D$_2$CO + H | 10              | $6 \times 10^{11}$            | $6 \times 10^{12}$ | 360         | 0.15 ± 0.04                  | 0.16 ± 0.01                  |
| H$_2$CO + H | 10              | $6 \times 10^{11}$            | $6 \times 10^{12}$ | 360         | 0.16 ± 0.04                  | 0.16 ± 0.01                  |
| D$_2$CO + H | 12              | $6 \times 10^{11}$            | $6 \times 10^{12}$ | 360         | 0.16 ± 0.04                  | 0.16 ± 0.01                  |
| H$_2$CO + H | 12              | $6 \times 10^{11}$            | $6 \times 10^{12}$ | 360         | 0.16 ± 0.04                  | 0.16 ± 0.01                  |
| D$_2$CO + H | 14              | $6 \times 10^{11}$            | $6 \times 10^{12}$ | 360         | 0.15 ± 0.04                  | 0.16 ± 0.01                  |
| H$_2$CO + H | 14              | $6 \times 10^{11}$            | $6 \times 10^{12}$ | 360         | 0.15 ± 0.04                  | 0.16 ± 0.01                  |
| D$_2$CO + H | 16              | $6 \times 10^{11}$            | $6 \times 10^{12}$ | 360         | 0.15 ± 0.04                  | 0.16 ± 0.01                  |
| H$_2$CO + H | 16              | $6 \times 10^{11}$            | $6 \times 10^{12}$ | 360         | 0.15 ± 0.04                  | 0.16 ± 0.01                  |
| D$_2$CO + H | 10              | $6 \times 10^{11}$            | $1.8 \times 10^{13}$ | 360         | 0.18 ± 0.03                  | 0.18 ± 0.01                  |
| H$_2$CO + H | 10              | $6 \times 10^{11}$            | $1.8 \times 10^{13}$ | 360         | 0.18 ± 0.03                  | 0.18 ± 0.01                  |
| H$_2$CO + D | 10              | $4 \times 10^{11}$            | $6 \times 10^{12}$ | 360         |                             |                             |

Note. The RAIRS abundances were calculated from the CO stretching bands ($\nu = 8$) of CHD$_2$OH and CH$_3$OH, and the TPD abundances from the $m/z = 32$ and $m/z = 34$ mass signals.
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