From spirals to lenticulars: Evidence from the rotation curves and mass models of three early-type galaxies

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ABSTRACT

Rotation curves have traditionally been difficult to trace for early-type galaxies (ETGs) because they often lack a high-density disk of cold gas as in late-type galaxies (LTGs). In this work, we derive rotation curves for three lenticular galaxies from the ATLAS$^{3D}$ survey, combining CO data in the inner parts with deep HI data in the outer regions, extending out to 10−20 effective radii. We also use Spitzer photometry at 3.6 μm to decompose the rotation curves into the contributions of baryons and dark matter (DM). We find that (1) the rotation-curve shapes of these ETGs are similar to those of LTGs of a similar mass and surface brightness; (2) the dynamically-inferred stellar mass-to-light ratios are small for quiescent ETGs but similar to those of star-forming LTGs; (3) the DM halos follow the same scaling relations with galaxy luminosity as those of LTGs; and (4) one galaxy (NGC 3626) is poorly fit by cuspy DM profiles, suggesting that DM cores may exist in high-mass galaxies too. Our results indicate that these lenticular galaxies have recently transitioned from LTGs to ETGs without altering their DM halo structure (e.g., via a major merger), and they could be faded spirals. We also confirm that ETGs follow the same radial acceleration relation as LTGs, reinforcing the notion that this is a universal law for all galaxy types.

Key words. galaxies: elliptical and lenticular, cD – galaxies: evolution – dark matter – galaxies: spiral – galaxies: kinematics and dynamics

1. Introduction

Extended rotation curves and mass models have been derived for hundreds of late-type galaxies (LTGs: spirals and irregulars), but only for a few early-type galaxies (ETGs: ellipticals and lenticulars) because they often lack high-density HI disks (e.g., Lelli et al. 2016). Traditionally, dynamical studies of ETGs have relied on different kinematic tracers. One approach is to combine stellar kinematics in the inner parts (e.g., Cappellari 2016) with discrete tracers such as planetary nebulae and globular clusters in the outer regions (e.g., Forbes et al. 2016). Alternatively, strong lensing (e.g., Gerhard 2013) and X-ray halos (e.g., Buote & Humphrey 2012) have also been used to trace the mass distribution of ETGs. These methods allowed one to study the dark matter (DM) content of ETGs, but a direct comparison with LTGs has been highly nontrivial since different methods require different assumptions (e.g., Janz et al. 2016; Lelli et al. 2017). For example, studies of stellar kinematics and discrete tracers rely on assumptions about the orbital anisotropy and the tracer population, studies of X-ray halos hold under the assumption of hydrostatic equilibrium, while strong lensing studies may depend on the lensing model and are generally limited to the inner galactic radii.

On the other hand, cold gas (both atomic and molecular) is a more straightforward tracer of the gravitational potential for a number of reasons: (1) the cold gas forms thin rotating disks in which the pressure support is negligible, so the rotation velocity directly traces the circular velocity of a test particle in the given gravitational potential; (2) noncircular motions are generally small, contrary to the ionized gas traced by Hα and [OIII] lines; and (3) the atomic gas, if present, can extend to large radii, well beyond the bright stellar component. However, analyzing the gas kinematics of ETGs as it is done for LTGs has long been challenging due to the low surface brightness of the gas emission (e.g., Weijmans et al. 2008; Struve et al. 2010). Deep, interferometric radio surveys have been able to detect and spatially resolve the HI emission in ETGs in group and field environments (Serra et al. 2012), showing that the atomic gas often forms regularly rotating disks extending out to large radii. Moreover, deep CO observations have revealed that many ETGs host inner disks of molecular gas (Alatalo et al. 2013; Davis et al. 2013). These key observational advances allow us to study the cold gas kinematics in ETGs in detail, providing new insights into the DM halos of different galaxy types.

In this work, we derive rotation curves and mass models of three ETGs from the ATLAS$^{3D}$ survey (Cappellari et al. 2011a): NGC 2824, NGC 3626, and UGC 6176. These are the only three galaxies in the ATLAS$^{3D}$ sample that have both an inner CO disk and an outer HI disk. We combine CO and HI kinematics to determine accurate rotation curves in a uniform way over the entire extent of the gaseous body. Some of the known properties of the galaxies are summarized in Table 1. They are morphologically classified as lenticulars (S0) and identified as “fast rotators” based on their observed stellar kinematics (Cappellari et al. 2011a). In particular, NGC 3626 is a well-studied object since it has a stellar component that counter-rotates with respect to the gaseous disk (Ciri et al. 1995; Jore et al. 1996; Haynes et al. 2000; Sil’chenko et al. 1995; Jore et al. 1996; Haynes et al. 2000; Sil’chenko et al. 2010; Mazzei et al. 2014).
also has a companion galaxy (UGC 6341, see Fig. 1) with significant HI emission (Serra et al. 2012), which we do not study here.

2. HI and CO moment maps

We analyze public HI and CO cubes from the Atlas$^{3D}$ website (Serra et al. 2012; Alatalo et al. 2013). For NGC 2824, we use a deeper HI cube that was kindly provided to us by Dr. Paolo Serra. The cube properties are summarized in Table 2. We perform our analysis using the GIPSY imaging processing system (GIPSY, van der Hulst et al. 1992) and$^{3D}$Barolo (3D-Based Analysis of Rotating Object via Line Observations, Di Teodoro & Fraternali 2015). We create masks that identify regions of gas emission using the “smooth and search” algorithm of$^{3D}$Barolo with the parameter \texttt{threshVelocity} set to 1 and \texttt{minPix} set to the beam size of each cube. We apply these masks to the full-resolution cubes to create intensity maps (moment-zero), summing all channels that contain line emission, and velocity maps (moment-one), estimating an intensity-weighted mean velocity. Velocity dispersion maps (moment-two) have little sensitivity to the endpoint of the beam (R. Avelino, private communication), which we use to model the HI and CO velocity fields. For all three galaxies, the HI distribution is extended beyond the stellar body, while the CO distribution is concentrated near the galaxy center. Both HI and CO velocity fields display regular rotation with no clear signs of strong noncircular motions. This is also evinced by position-velocity (PV) diagrams extracted along the kinematic major axis of the disk, which are shown in Fig. 2. The morphological position angle (PA) from the intensity maps agrees with the kinematic PA from the velocity maps. It also agrees with the orientation of the stellar bodies for NGC 2824 and UGC 6176. For NGC 3626, instead, the PA suggested by the elongation of the stellar disk differs from the average HI major axis, pointing to a warped gas disk.

3. Derivation of rotation curves

We derive rotation curves by fitting the velocity fields with a tilted-ring model to obtain initial estimates, which are subsequently refined by building kinematic 3D disk models (e.g., Lelli et al. 2012a,b). We also verified our results using the software$^{3D}$Barolo (Di Teodoro & Fraternali 2015), which directly fits the 3D cube bypassing the 2D velocity maps.

In the tilted-ring fit, each ring is described by six free parameters: center (x⊙, y⊙), systemic velocity (v⊙), position angle (PA), inclination (i), and rotation velocity (vrot). The width of the rings is chosen to be either equal to the width of the beam major axis, or \( \sqrt{(b_{\text{maj}}/b_{\text{min}})} \) if the beam is strongly elongated and not aligned with the PA of the galaxy. Here, \( b_{\text{maj}} \) and \( b_{\text{min}} \) are the major and minor axes of the beam, respectively. The number of rings is determined such that the major axis of the galaxy is fully covered. The pixels in the velocity maps are weighted using a cosine function to give more weight to points close to the major axis, which provide most information regarding the rotation velocity. We simultaneously fit both halves of the galaxy, neglecting possible asymmetries between the different sides. For NGC 3626, however, the CO data do not cover the full gas emission at approaching velocities due to the limited bandwidth of the observations, so we fit only the receding side.

We perform a series of iterations using the GIPSY task \texttt{rotcur} (Begeman 1987). In the first iteration all the parameters are left free, while in the second iteration we fix \( x_0 \), \( y_0 \), and \( v_0 \) using the mean values across the rings. The centers and systemic velocities from the CO and HI data are consistent with each other, as well as with optical determinations from the Nasa Extragalactic Database. In the next two iterations, the PA and inclination are fixed in arbitrary order. We have checked that the order does not change the resulting rotation curve within the errors. For NGC 2824 and UGC 6176, the CO and HI disks show no strong indications of warps, so i and PA are kept constant with radius. For both galaxies, however, the inner CO inclination significantly differs from the outer HI inclination. For NGC 3626, we model an HI warp by varying both PA and i with radius. The last run of \texttt{rotcur} is performed leaving only \( v_{\text{rot}} \) as a free parameter.

To check the resulting rotation curves and correct for beam smearing effects, we construct model cubes using the \texttt{rotcur} results and the observed gas density profiles as inputs. We derive azimuthally-averaged HI and CO surface brightness profiles using the GIPSY task \texttt{ellint}, adopting the best-fit geometric parameters. The CO surface brightnesses are converted into Hz surface densities adopting the Milky-Way conversion factor \( X_{\text{CO}} = 2 \times 10^{20} \text{cm}^{-2} (K \text{km} \text{s}^{-1})^{-1} \) (Bolatto et al. 2013). We create model cubes with the GIPSY task \texttt{galmod} for the HI data and a custom-built task for the CO data. Since the latter task does no automatic conversion from Hz surface density to CO surface brightness, we renormalize the total flux of the CO model to be equal to the observed flux. We assume that the gas disk has a velocity dispersion of 10 km s$^{-1}$ and a thickness of 100 pc.

We compare PV diagrams and channel maps from the model cubes with the observed ones. A disagreement between models

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Table 1. Galaxy sample.

|       | NGC 2824 | NGC 3626 | UGC 6176 |
|-------|----------|----------|----------|
| RA    | 09 19 02.23 | 11 20 03.79 | 11 07 24.66 |
| Dec   | +26 16 12.00 | +18 21 24.45 | +21 39 25.87 |
| Morph. type | S0 | SA0 | SB0 |
| D [Mpc] | 39.6 | 22.9 | 39.3 |
| M_K [mag] | −22.87 | −23.65 | −22.56 |
| log M_HI/M_⊙ | 8.93 | 9.08 | 9.00 |
| log M_CO/M_⊙ | 8.89 | 8.46 | 8.74 |
| log L_{3.6}/L_⊙ | 10.6 | 10.9 | 10.5 |
| R_{eff} [kpc] | 1.09 | 2.43 | 1.18 |
| Σ_{eff} [L_⊙/pc²] | 5210 | 2018 | 3318 |

Notes. The galaxy coordinates are from Cappellari et al. (2011a). The morphological type is from de Vaucouleurs et al. (1991). NGC 3626 hosts a bar and has been reclassified as SAB0/a by Buta et al. (2015). The galaxy distance D is from Lelli et al. (2017). The K-band absolute magnitude M_K (corrected for foreground galactic extinction), the atomic gas mass M_HI, and the molecular gas mass M_CO are from Cappellari et al. (2011a), Serra et al. (2012) and Alatalo et al. (2013), respectively, and converted to the adopted distances. The HI mass of NGC 2824 is estimated in this work using deeper observations. The luminosity at 3.6μm, the effective radius R_{eff} encompassing half the total luminosity, and the effective surface brightness Σ_{eff} are from Lelli et al. (2017).
Fig. 1. Moment maps for NGC 2824 (left), NGC 3626 (middle), and UGC 6176 (right). The top section shows the HI data, while the bottom section shows the CO data. In each section, the top row shows the Spitzer image at 3.6 µm overlaid with the gas intensity maps, while the bottom row shows the velocity field. The box in the HI velocity field depicts the size of the CO frame. The cross corresponds to the galaxy center; the dashed line indicates the kinematic major axis; the physical scale is illustrated in the bottom-right corner; the beam shape is shown in the bottom-left corner. The HI and CO density contours are multiples of the pseudo-3σ value (0.5, 1, 2, 4, …) where 3σ map is given in Table 2.

and observations indicates that the input density profiles and/or rotation curves are affected by beam smearing. Beam smearing tends to flatten the surface density profile and cause an underestimation of the rotation velocity, especially in the innermost regions. We manually adjust individual points in the inner parts of the rotation curves and/or density profiles to minimize the residuals. In some cases, the PA and inclination angle are also manually adjusted by a few degrees to find the best compromise between the kinematic estimates from rotcur and the observed intensity maps. The best-fit 3D models are illustrated in Fig. 2 and presented in a more extensive way in Appendix A. The adopted geometric parameters are summarized in Tables 3 and 4 for the warp model of NGC 3626. The uncertainties are estimated by running 3DBarolo with all the parameters fixed to the final value and only one parameter of interest left free. The final geometric parameters and rotation curves are shown in Fig. 3.

In general, 3DBarolo returns similar rotation curves as our manual procedure with some exceptions. For NGC 2824, the CO data have too low spatial resolutions and signal-to-noise ratio
to ensure a good 3D Barolo fit, while the HI data shows a strong absorption feature at small radii that 3D Barolo cannot model. For NGC 3626, the missing CO data at approaching velocities can- not be automatically modeled by 3D Barolo, whereas our manual procedure can take this shortcoming into account by focusing on the receding portion of the PV diagram. For these reasons, we prefer our manual procedure with respect to fully automatic fits with 3D Barolo. The rotation curves are often observed in massive spiral galaxies (see, e.g., de Grijs 1998). Luckily, the 3D geometry of baryons has a minor effect on the final mass models: the maximum variation in the peak circular velocities is about 20% for axial ratios between zero (razor-thin disk) and one (sphere); see, for example, Fig. 3 in Noordermeer (2008). This effect is smaller than the current uncertainties in the stellar mass-to-light ratios from stellar population models (e.g., Schombert et al. 2019). Having computed the gravitational field of the individual components in the disk mid-plane, the baryonic rotation curve reads

\[
v_{\text{curv}}(R) = v_d v_1^2(R) + v_b v_2^2(R) + 1.33 \left[ v_{\text{HI}}^2(R) + v_{\text{H}_2}^2(R) \right],
\]

where \(v_1\) and \(v_2\) account for the stellar contribution (disk and bulge, respectively), while \(v_{\text{HI}}\) and \(v_{\text{H}_2}\) account for the gas contribution (atomic and molecular, respectively). The molecular contribution is computed assuming the standard conversion factor \(X_{\text{CO}} = 2 \times 10^{20} \text{cm}^{-2} (\text{K} \text{km s}^{-1})^{-1}\) of the Milky Way. \(\Upsilon\) denotes the mass-to-light ratio for the stellar disk and the bulge, indicated by the subscripts \(d\) and \(b\), respectively. The 1.33 factor in front of the gaseous components takes the mass contribution of primordial helium into account (Lelli et al. 2016; but see also McGaugh et al. 2020). Table 5 provides the circular velocities and the surface densities of the individual components at the same radii at which rotation velocities are measured.

### 4. Mass models

We derive detailed mass models using Spitzer 3.6\,μm images to estimate the stellar gravitational field, together with HI and CO maps to estimate the gas gravitational field (Fig. 1). In Sect. 5, we use these mass models and the observed rotation curves to study the DM halos of our three lenticular galaxies.

#### 4.1. Stellar and gaseous contributions

The SPARC database (Spitzer Photometry and Accurate Rotation Curves, Lelli et al. 2016) provides [3.6] luminosity profiles and nonparametric bulge-disk decompositions for several ETGs in Atlas3D including our three S0s (see Lelli et al. 2017). Given the intrinsic ambiguity of bulge-disk decompositions for lenticular galaxies, we explored different decompositions’ scheme to estimate their impact on the final mass models. After several trials, we decided to use a single, disk-like component for both NGC 2824 and NGC 3626 because the bulge contribution is small and uncertain. This is a conservative choice that avoids adding unnecessary free parameters to the mass model. All the observed features in the surface brightness profiles are preserved in calculating the stellar gravitational contribution, including the inner light concentrations. For NGC 2824 and NGC 3626, we simply assume that these light concentrations lie in a flattened structure with a similar mass-to-light ratio as the stellar disk, as it may be the case for a pseudo-bulge (e.g., Kormendy & Freeman 2004). For UGC 6176, we use the same bulge-disk decomposition as in Lelli et al. (2017).

We calculate the gravitational contribution of each component using the GIPSY task rotmod. We assume a spherical bulge (when present) and thick disks with a vertical exponential profile. The scale-heights of the HI and CO disks are set to 100 pc, while that of the stellar disk \((h_z)\) is assumed to correlate with the exponential scale-length \((h_z)\) as \(h_z = 0.196 h_{10}^{0.633}\) (Bershady et al. 2010). This relation has been derived for edge- on spiral galaxies and may be potentially inappropriate for lenticular galaxies (see, e.g., de Grijs 1998). Luckily, the 3D geometry of baryons has a minor effect on the final mass models: the maximum variation in the peak circular velocities is about 20% for axial ratios between zero (razor-thin disk) and one (sphere); see, for example, Fig. 3 in Noordermeer (2008). This effect is smaller than the current uncertainties in the stellar mass-to-light ratios from stellar population models (e.g., Schombert et al. 2019). Having computed the gravitational field of the individual components in the disk mid-plane, the baryonic rotation curve reads

\[
v_{\text{curv}}(R) = \Upsilon_d v_1^2(R) + \Upsilon_b v_2^2(R) + 1.33 \left[ v_{\text{HI}}^2(R) + v_{\text{H}_2}^2(R) \right],
\]

where \(v_1\) and \(v_2\) account for the stellar contribution (disk and bulge, respectively), while \(v_{\text{HI}}\) and \(v_{\text{H}_2}\) account for the gas contribution (atomic and molecular, respectively). The molecular contribution is computed assuming the standard conversion factor \(X_{\text{CO}} = 2 \times 10^{20} \text{cm}^{-2} (\text{K} \text{km s}^{-1})^{-1}\) of the Milky Way. \(\Upsilon\) denotes the mass-to-light ratio for the stellar disk and the bulge, indicated by the subscripts \(d\) and \(b\), respectively. The 1.33 factor in front of the gaseous components takes the mass contribution of primordial helium into account (Lelli et al. 2016; but see also McGaugh et al. 2020). Table 5 provides the circular velocities and the surface densities of the individual components at the same radii at which rotation velocities are measured.

#### 4.2. Maximum-disk mass models

The baryonic rotation curve depends on the adopted stellar mass-to-light ratios (Eq. (1)). Here we adjust their values to achieve a maximum-disk fit (van Albada & Sancisi 1986; Starkman et al. 2018), ensuring that the inner part of the observed rotation curve is fully explained by luminous matter. Our fits can be seen in Fig. 4 along with the maximal values of \(\Upsilon_b\) and \(\Upsilon_d\). Surprisingly, the maximum mass-to-light ratios are relatively small for red and dead ETGs, which are expected to be in the range 0.8–1.0 \(M_\odot/L_\odot\) (Schombert & McGaugh 2014), but are near the high-end boundary of the expectations for star-forming LGTs (Schombert et al. 2019). This is further discussed in Sect. 6. The
Fig. 2. Major-axis PV diagrams from the observed HI cube (top row) and the CO one (bottom row), overlaid with the projected rotation curve (white squares) and model PV diagrams (blue contour corresponding to 1.5σ from Table 2).

Table 3. Geometric parameters of CO and HI disks.

|          | NGC 2824 | NGC 3626 | UGC 6176 |
|----------|----------|----------|----------|
| RA (hms) | 09 19 2.23 | 11 20 4.15 | 11 07 24.70 |
| Dec (dms) | +26 16 12.0 | +18 21 24.5 | +21 39 25.9 |
| $V_{sys}$ (km s$^{-1}$) | 2725 ± 4 | 1468 ± 5 | 2680 ± 3 |
| PA (°)    | 160 ± 9 | -- | 33 ± 7 |
| $i_{CO}$ (°) | 50 ± 2 | -- | 85 ± 1 |
| $i_{HI}$ (°) | 70 ± 2 | -- | 58 ± 1 |

Notes. The coordinates of the kinematic center (RA and Dec) are given in J2000. We adopt the same systemic velocity ($V_{sys}$) and position angle (PA) for CO and HI disks, but we use different inclinations ($i$). The PA and $i$ of NGC 3626 are radius-dependent and can be found in Table 4.

Table 4. Warp model for NGC 3626.

|          | R ("') | R (kpc) | $i$ (°) | PA (°) |
|----------|--------|---------|--------|--------|
| 15       | 1.67   | 50      | 174    |
| 45       | 4.50   | 50      | 174    |
| 75       | 8.33   | 53      | 178    |
| 105      | 11.66  | 57      | 180    |
| 135      | 14.99  | 60      | 180    |
| 165      | 18.32  | 60      | 180    |
| 195      | 21.65  | 60      | 180    |
| 225      | 24.99  | 60      | 180    |

Notes. Sampled radii (R) in both arcsec and kpc, adopted inclination ($i$), and position angle (PA) for the warped disk of NGC 3626.

baryonic rotation curves do not agree with the observed ones at large radii. This indicates the presence of the DM effect in the galaxies, as we investigate in the next Section.

5. Dark matter halo fits

To model the discrepancy between the observed and baryonic rotation curves at large radii (Fig. 4), we add a DM halo to the mass models. We fit three different halo profiles: NFW (Navarro et al. 1996), Einasto (Einasto 1965), and DC14 (Di Cintio et al. 2014). While the NFW and Einasto profiles are suggested by DM-only cosmological simulations (e.g., Dutton & Macciò 2014), the DC14 profile is derived from hydrodynamical simulations of galaxy formation, and is taking into account the effects of the stellar feedback on the inner DM halo.

Rotation-curve fits with the DC14 profile are described in Katz et al. (2017) and Li et al. (2019, 2020). In short, the shape of the DC14 profile depends on the quantity $X = \log(M_\star/M_{200})$ where $M_\star$ is the present-day stellar mass and $M_{200}$ is the halo mass, defined as the enclosed mass at which the halo density equals 200 times the critical density of the Universe. For $-4 < X < -2$, the DC14 profile is shallower than the NFW profile due to the formation of central DM cores by stellar feedback. For higher and lower values of $X$, the DC14 reduces to the NFW profile because the supernovae energy is too small to overcome the galaxy potential well and drive a significant mass redistribution. Fox $X > -1.3$, however, the feedback from active galactic nuclei (AGN), which is not included in the simulations of Di Cintio et al. (2014), is expected to become...
important, making the DC14 profile potentially inaccurate. After performing the fits we check that $X < -1.3$ for all galaxies, so the use of the DC14 profile is justified. Specifically, we find $-1.37$ for NGC 2824, $-1.99$ for NGC 3626, and $-1.44$ for UGC 6176. These stellar-to-halo mass ratios are rather high, so the DC14 results are expected to be similar to the NFW and Einasto ones.

5.1. ΛCDM priors

Rotation-curve fits with DM halos are subject to a number of degeneracies due to the significant number of fitting parameters. The availability of Spitzer [3.6] photometry greatly alleviates the so-called disk-halo degeneracy (van Albada et al. 1985), but additional “priors” in a Bayesian context are needed to fully break the problem degeneracies. Thus, it is sensible to constrain the free parameters to domains that are in agreement with cosmological expectations. Specifically, we impose two relations expected from ΛCDM cosmological simulations, following the same procedures as in Li et al. (2019, 2020).

Dutton & Macciò (2014) used cosmological DM-only simulations to study the relation between the mass ($M_{200}$) and concentration ($C_{200}$) of the halo. This is given by

$$ \log c_{200} = a - b \log \left( \frac{M_{200}}{10^{12} \, h^{-1} \, M_\odot} \right), \quad (2) $$

with an intrinsic scatter of 0.11 dex. Here $a$ and $b$ are cosmology-dependent parameters, while $h = H_0/100$ where $H_0$ is the Hubble constant. We assume $H_0 = 73 \, \text{km} \, \text{s}^{-1} \, \text{Mpc}^{-1}$ to be consistent with the local SPARC distance scale (Lelli et al. 2016).

Moster et al. (2013) derived a stellar-to-halo mass relation by matching the number of observed galaxies in the local Universe to the predicted number of DM halos from $N$-body simulations. The functional form of the stellar-to-halo-mass (SHM) relation is assumed to be

$$ \frac{M_*}{M_{200}} = 2N \left[ \left( \frac{M_{200}}{M_1} \right)^{\beta} + \left( \frac{M_{200}}{M_1} \right)^{\gamma} \right]^{-1}, \quad (3) $$

where $\log(M_1) = 11.59$, $N = 0.0351$, $\beta = 1.376$ and $\gamma = 0.608$. Without imposing these two relationships, the parameters of the DM halos are poorly constrained (see Appendix B), so in the following we mostly focus on rotation-curve fits with log-normal ΛCDM priors.

5.2. Fitting procedure

To fit the DM halos, we follow Li et al. (2019, 2020) and use their Markov-chain-Monte-Carlo (MCMC) code, which maps the posterior distributions of the fitting parameters using the python package emcee (Foreman-Mackey et al. 2013). The code is supplied the observed rotation curve and the circular velocities of the individual baryonic components. The free parameters are the halo mass and halo concentration (plus the shape parameter $\alpha_x$ for the Einasto profile), the stellar mass-to-light ratios, the disk inclination ($i$), and the galaxy distance ($D$). For convenience, we perform the fits using $V_{200} = [10G H_0 M_{200}]^{1/3}$ (McGaugh 2012), where $G$ is the gravitational constant. We assume Gaussian priors on $i$ and $D$ centered at the estimated values and with standard deviations given by the estimated uncertainties. Since our galaxies show radial variations of $i$, we assume that the inclination uncertainty is equal to half of the total variation. We adopt log-normal priors on the mass-to-light ratios, centered at 0.5 for the disk and 0.7 for the bulge with a standard deviation of 0.1 dex.

For the halo parameters, we use loose boundaries to keep them in a physically sensible range: $10 < V_{200} < 500 \, \text{km} \, \text{s}^{-1}$, $0 < c_{200} < 100$, and for the Einasto profile $0 < \alpha_x < 2$. When we apply the ΛCDM priors, we need to rescale the concentration of the DC14 profile since it is defined in a slightly different

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Fig. 3. Position angle (top), inclination (middle), and final rotation curves (bottom) from HI (green points) and CO (blue triangles) observations. Open circles show omitted points in the combined rotation curves due to the overlap of CO and HI data.
### Table 5. Rotation curves and mass models for three lenticular galaxies.

| Galaxy  | $R$ [kpc] | $v_{\text{obs}}$ [km s$^{-1}$] | $e_{v_{\text{obs}}}$ | $v_{\text{HI}}$ [km s$^{-1}$] | $v_{\text{disk}}$ [km s$^{-1}$] | $v_{\text{bulge}}$ | $\Sigma_{\text{HI}}$ [L$\odot$ pc$^{-2}$] | $\Sigma_{\text{HI}}$ [L$\odot$ pc$^{-2}$] | $\Sigma_{\text{HI}}$ [L$\odot$ pc$^{-2}$] | $\Sigma_{\text{HI}}$ [L$\odot$ pc$^{-2}$] |
|---------|-----------|-------------------------------|------------------|-----------------|-----------------|-----------------|-----------------|-----------------|-----------------|-----------------|
| NGC 2824| 0.38      | 140                           | 29               | 17.82           | 0.95            | 222.95          | –               | 120.89          | 2.61            | 8422.37         |
|         | 1.15      | 194                           | 29               | 39.08           | 2.83            | 243.31          | –               | 60.77           | 2.54            | 1238.20         |
|         | 1.92      | 211                           | 29               | 44.22           | 4.48            | 234.68          | –               | 18.87           | 2.46            | 506.32          |
| NGC 3626| 2.88      | 160                           | 10               | 38.55           | 6.39            | 209.06          | –               | 3.96            | 2.37            | 139.56          |
|         | 4.81      | 160                           | 11               | 28.95           | 10.31           | 169.19          | –               | 0.19            | 2.25            | 59.68           |
|         | 6.73      | 160                           | 8                | 23.76           | 14.80           | 151.28          | –               | 0.01            | 1.90            | 17.84           |
|         | 8.65      | 160                           | 10               | 20.70           | 17.82           | 132.94          | –               | 0.00            | 1.28            | 10.12           |
|         | 0.22      | 147                           | 15               | 7.48            | -0.22           | 199.14          | –               | 63.93           | 2.04            | 23094.69        |
|         | 0.67      | 182                           | 15               | 19.16           | 1.75            | 265.53          | –               | 49.70           | 1.97            | 3566.60         |
|         | 1.11      | 214                           | 15               | 28.68           | 2.78            | 258.78          | –               | 34.04           | 1.90            | 1747.42         |
|         | 1.55      | 220                           | 30               | 32.82           | 3.70            | 251.56          | –               | 13.32           | 1.82            | 1233.41         |
|         | 2.00      | 237                           | 30               | 29.53           | 4.55            | 254.81          | –               | 0.74            | 1.75            | 1048.37         |
|         | 4.99      | 197                           | 15               | 16.26           | 8.79            | 239.53          | –               | 0               | 1.26            | 157.55          |
| UGC 06176| 8.32      | 157                           | 12               | 9.44            | 10.36           | 199.06          | –               | 0               | 0.87            | 36.76           |
|         | 11.65     | 158                           | 11               | 10.50           | 11.29           | 173.70          | –               | 0               | 0.92            | 13.83           |
|         | 14.98     | 169                           | 7                | 9.25            | 15.36           | 154.07          | –               | 0               | 0.92            | 3.18            |
|         | 18.31     | 180                           | 9                | 8.37            | 20.41           | 138.64          | –               | 0               | 0.49            | 0               |
|         | 21.64     | 191                           | 12               | 7.70            | 18.97           | 124.74          | –               | 0               | 0.21            | 0               |
|         | 24.97     | 198                           | 11               | 7.15            | 17.64           | 115.31          | –               | 0               | 0.09            | 0               |
|         | 0.38      | 120                           | 14               | 30.50           | 0.74            | 20.39           | 125.74          | 229.04          | 1.83            | 390.16          |
|         | 3.81      | 137                           | 8                | 27.18           | 6.53            | 87.10           | 154.91          | 0.06            | 1.55            | 175.01          |
|         | 11.43     | 137                           | 8                | 15.15           | 14.76           | 56.18           | 50.28           | 0               | 0.87            | 1.92            |
|         | 19.05     | 137                           | 8                | 11.77           | 17.81           | 41.33           | 38.95           | 0               | 0.29            | 0               |

**Notes.** Sampled radii ($R$), observed rotation velocities ($v_{\text{obs}}$), and relative errors ($e_{v_{\text{obs}}}$) from CO and HI datacubes. Contributions to the circular velocity from the molecular gas disk ($v_{\text{HI}}$), atomic gas disk ($v_{\text{HI}}$), stellar disk ($v_{\text{disk}}$), and -if present- stellar bulge ($v_{\text{bulge}}$), together with the corresponding surface densities at $R$. For convenience, the velocity contributions and surface density profiles of the stellar disk and the bulge are normalized to a mass-to-light ratio of 1 at 3.6 $\mu$m.
Fig. 4. Top panels: surface density profiles for the H₂ disk (blue circles), HI disk (green upward triangles), stellar disk (yellow downward triangles), and stellar bulge (red diamonds) when necessary. Open circles indicate that the surface density is extrapolated from an exponential fit. Bottom panels: observed rotation curve (dots with ±1σ error bars) is fit with a maximum-disk mass models (black solid line with ±1σ error band). The baryonic rotation curve is decomposed into the contributions of molecular gas (blue dashed line), atomic gas (green dash-dotted line), stellar disk (yellow solid line), and stellar bulge (red dotted line) when necessary. The maximum mass-to-light ratios are given in the top right corner.

way from that of the NFW and Einasto profiles (Di Cintio et al. 2014). Following Katz et al. (2017), we adopt

$$\rho_{\text{rescaled}} = c_{\text{DC14}} \times [1 + 10^{-5} \exp(3.4X + 15.3)]^{-1}. \quad (4)$$

To impose the mass-concentration relation, we adopt $a = 0.830$ and $b = -0.098$ for the NFW and DC14 profiles, assuming a WMAP5 cosmology. This is the best compromise between the WMAP3 cosmology assumed in the simulations of Di Cintio et al. (2014) and the WMAP7 cosmology assumed in the SHM relation of Moster et al. (2013). However, as stated in Li et al. (2019), the difference induced by different cosmological parameters is nearly negligible and mostly plays a role in the final distance estimates. For the Einasto profile, we adopt $a = 0.977$ and $b = -0.130$ assuming a Planck cosmology since this is the only available calibration. We also impose an additional constraint on $\alpha_e$ describing the dependence of this parameter on halo mass. According to Dutton & Macciò (2014), this is given by

$$\alpha_e = 0.0095 \nu^2 + 0.155, \quad (5)$$

with $\log \nu = -0.11 + 0.146m + 0.0138m^2 + 0.00123m^3$, $m = \log(M_{\text{halo}}/10^{12} h^{-1} M_\odot)$, and a vertical scatter of 0.16 dex.

For each fit, we perform 2000 burn-in steps and 1000 iterations after that. Then, we re-initialize the walkers around the maximum-probability values and perform 1000 more iterations. We have visually checked that the chains have converged. In cases where the posterior distributions have complex shapes, we have checked that changing the number of the burn-in steps

1 There is a typo in Eq. (3) of Katz et al. (2017) since the factor $10^{-5}$ appears inside the exponential. The equation written here is the correct one, which was used in the analysis of Katz et al. (2017).
to the right of the SHM relation because its rotation curve is pushing the fit toward low $M_\star/M_{200}$ values in order to have a less cuspy DM profile. Figure 7 (bottom panels) shows that the halo parameters are also consistent with the imposed mass-concentration relation (Dutton & Macciò 2014) within the large uncertainties.

6. Discussion

6.1. Radial acceleration relation

Late type galaxies follow a tight radial acceleration relation (RAR; McGaugh 2016; Lelli et al. 2017): the observed centripetal acceleration at each radius correlates with that predicted by the baryonic components. This implies that, to a first approximation, the rotation-curve shape can be predicted by the baryon distribution, and vice versa. The scatter around the RAR is remarkably small for astronomical standards and is largely dominated by observational scatter, leaving little space for intrinsic scatter in the relation (Li et al. 2018, but see also Stone & Courteau 2019). In addition to the LTGs from the SPARC sample, Lelli et al. (2017) also considered 25 ETGs and 62 dwarf spheroidals, finding that they follow the same relation as LTGs (with the possible exception of ultra-faint dwarfs for which the observational situation is more uncertain). The RAR might represent a universal law of nature with severe implications on cosmological models (see Lelli et al. 2017 for a discussion).

Early type galaxies are the least represented galaxy type in the RAR. The sample of Lelli et al. (2017) included nine so-called slow rotators and 16 fast rotators (Cappellari et al. 2011b). The slow rotators are generally giant boxy ellipticals, in which the observed centripetal acceleration can be measured using X-ray halos in hydrostatic equilibrium (Buote & Humphrey 2012). The fast rotators comprise both disky ellipticals and lenticular galaxies: their inner circular velocities were estimated fitting Jeans axisymmetric models to the stellar kinematics (Cappellari et al. 2011b), while the outer circular velocities were measured using HI observations (den Heijer et al. 2015). Therefore, the mass models of these fast rotators were limited to two measured points only. Having derived full, extended rotation curves for rotating ETGs, we can more rigorously investigate their location on the RAR.

Following Lelli et al. (2017), we calculate the observed centripetal acceleration as $\phi_{\text{obs}} = v_{\text{obs}}^2(R)/R$ and the one predicted from the baryon distribution as $\phi_{\text{bar}} = v_{\text{bar}}^2(R)/R$ with $\Upsilon_d = 0.5$ and $\Upsilon_b = 0.7$. Figure 8 shows that our ETGs lie on top of the
RAR derived for LTGs within 1σ. Although our sample is too small to make statistically robust conclusions, this result suggests that galaxies follow the same RAR independently of their morphology and evolutionary stage. Thus, the rotation curves of ETGs can be predicted from the baryonic mass distribution using the same relation as for LTGs. This points to the possible universality of the RAR, in line with the predictions of Milgromian dynamics (MOND, Milgrom 1983). A bigger ETG sample with both CO and HI data is needed to reinforce this result.

### 6.2. Evolution of early-type galaxies

The dichotomy of early-type galaxies, dividing them into slow and fast rotators (e.g., Emsellem et al. 2007), is well established but the formation and evolution of the two galaxy types is subject of debate. The galaxies studied in this paper are rotating lenticulars, so we focus on their formation and evolution.

Several authors (e.g., Kormendy & Bender 2012; Laurikainen et al. 2010) proposed that lenticulars evolve from spirals through secular quenching of star formation, possibly

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**Table 6.** Maximum-probability values with ΛCDM constraints for the fit parameters.

| Galaxy   | Halo profile | $D$ [Mpc] | $i$ [°] | $T_d$ | $T_b$ | $V_{200}$ [km s$^{-1}$] | $C_{200}$ | $\alpha_e$ |
|----------|--------------|-----------|---------|-------|-------|-------------------------|-----------|-----------|
| NGC 2824 | NFW          | 43 ± 7    | 59 ± 9  | 0.47 ± 0.09 | –     | 130 ± 28 | 8 ± 2 | –         |
|          | Einasto      | 37 ± 8    | 59 ± 9  | 0.47 ± 0.10 | –     | 114 ± 21 | 11 ± 3 | 0.18 ± 0.07 |
|          | DC14         | 41 ± 7    | 59 ± 8  | 0.46 ± 0.09 | –     | 113 ± 28 | 11 ± 6 | –         |
| NGC 3626 | NFW          | 23 ± 2    | 62 ± 8  | 0.54 ± 0.10 | –     | 169 ± 33 | 5 ± 1 | –         |
|          | Einasto      | 22 ± 2    | 61 ± 8  | 0.56 ± 0.11 | –     | 160 ± 56 | 5 ± 2 | 0.3 ± 0.1 |
|          | DC14         | 25 ± 2    | 64 ± 7  | 0.52 ± 0.08 | –     | 228 ± 44 | 7 ± 1 | –         |
| UGC 6176 | NFW          | 40 ± 5    | 56 ± 2  | 0.49 ± 0.10 | 0.69 ± 0.14 | 110 ± 10 | 8 ± 2 | –         |
|          | Einasto      | 35 ± 5    | 56 ± 2  | 0.48 ± 0.10 | 0.70 ± 0.15 | 102 ± 9  | 10 ± 3 | 0.17 ± 0.05 |
|          | DC14         | 40 ± 6    | 56 ± 2  | 0.47 ± 0.10 | 0.69 ± 0.14 | 109 ± 26 | 10 ± 4 | –         |
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\[ \log_{10}\left( \frac{M_{\text{halo}}}{M_{\odot}} \right) \]

Einasto

\[ \log_{10}\left( \frac{M_{\star}}{M_{\odot}} \right) \]

NGC3626

NGC2824

UGC6176

NFW

\[ \log_{10}\left( \frac{M_{\text{halo}}}{M_{\odot}} \right) \]

DC14

\[ \log_{10}\left( \frac{M_{\star}}{M_{\odot}} \right) \]

\[ \log_{10}\left( \frac{M_{200}}{M_{\odot}} \right) \]

\[ \log_{10}\left( c_{200} \right) \]

Fig. 7. Stellar-to-halo-mass relation (top) and halo mass-concentration relation (bottom). The black solid line and the gray shaded areas show the relations expected in \( \Lambda \)CDM together with their intrinsic scatter (Moster et al. 2013; Dutton & Macciò 2014). The solid points with ±1 sigma errorbars show the maximum-probability values from rotation-curve fits with \( \Lambda \)CDM priors.

Fig. 8. RAR for our ETGs using \( Y_d = 0.5 \) and \( Y_b = 0.7 \) as in Lelli et al. (2017). Black and dashed lines correspond, respectively, to the mean and standard deviation of the RAR for LTGs in the SPARC database. The dotted line shows the 1:1 relation.

Driven by gas stripping and/or starvation (i.e., lack of gas accretion). These conclusions are motivated by the fact that the stellar disks and bulges of S0s follow the same photometric scaling relations as those of spirals. The fading scenario is backed from comparative population studies, showing that the fraction of LTGs decreases as a function of environmental density with nearly the same slope as the fraction of rotating ETGs increases (Cappellari et al. 2011b). Numerical simulations, however, show that lenticulars formed via major mergers may have coupled bulge-disk components that follow the same scaling relations as spirals (Querejeta et al. 2015), thus photometric data alone may...
Fig. 10. DM halo parameters versus [3.6] luminosity for our three ETGs and the LTGs in the SPARC sample (gray dots from Li et al. 2019): halo scale radius (top panels), characteristic volume density (middle panels), and their product (bottom panels).

not be sufficient to distinguish between secular fading and major mergers. In this evolutionary context, kinematical and dynamical studies acquire primary importance.

Rizzo et al. (2018) performed kinematic bulge-disk decompositions of unbarred S0s, separating the light contributions from bulge and disk at each spatial position of the optical spectrum. They found that the stellar disks of S0s have the same specific angular momentum as those of spirals, concluding that major mergers are not the dominant formation mechanisms for S0s in the group and field environments. Diametrically different conclusions are drawn by Falcón-Barroso et al. (2019), which, however, did not perform a kinematic bulge-disk decomposition of optical spectra but included bulges in their estimate of the specific stellar angular momentum. The results of our work, based on gas kinematics that is exempt from highly nontrivial spectral decompositions, give further support to the scenario in
which S0s are forming from aging spiral galaxies in a fluent transition.

Figure 9 compares the rotation curves of ETGs with those of LTGs from SPARC (Lelli et al. 2016). The rotation-curve shapes of S0s are similar to spirals of comparable effective surface brightness and flat rotation velocity, hence they are of comparable baryonic mass (stars plus gas) since HI-rich ETGs follow the same baryonic Tully-Fisher relation as LTGs (den Heijer et al. 2015). This implies that the shape of the total gravitational potential is similar in spirals and lenticulars. Moreover, the amount of specific angular momentum in the gas disk of S0s is clearly similar to that of spirals, in analogy to the specific angular momentum in the stellar disk (Rizzo et al. 2018). Finally, our values for the mass-to-light ratios (0.5–0.7 $M_\odot/L_\odot$ at 3.6 $\mu$m) are lower than expected for old and passive ETGs with exponentially declining star-formation histories (0.8–1.0 $M_\odot/L_\odot$, see Fig. 11 in Schombert & McGaugh 2014), but are in the typical range of star-forming LTGs (Schombert et al. 2019). This suggests that the star-formation activity in our three ETGs must have dropped (if not stopped) only recently, which allows us to classify them as faded spirals.

Our findings are consistent with those of Yıldız et al. (2017), who used a sample of HI-rich and HI-poor ETGs from the ATLAS$^3D$ survey to study their star-formation properties. They compared the UV colors of the two samples and found that HI-rich ETGs are bluer than HI-poor ETGs. The most gas-rich ETGs are as blue as the outer parts of typical LTGs. Yıldız et al. (2017) concluded that gas-rich ETGs host young stellar populations in their outskirts. They also found evidence of low star formation efficiency in these regions. Although our galaxies are not part of Yıldız’s sample, their HI mass places them among the gas-rich ETGs (NGC 2824 was not considered by Yıldız et al. (2017) due to the weak HI detection in the Atlas$^3D$ data, while NGC 3626 and UGC 6176 had no available UV observations). NGC 3626 likely has some ongoing star formation, possibly triggered by a recent interaction with its companion galaxy, as suggested by the stellar and gas rings visible in its outer disk.

Further support for a faded-spiral scenario comes from comparing our derived DM halo parameters to the ones for the LTGs in the SPARC sample (Li et al. 2019). This is shown in Fig. 10, where the halo scale length and characteristic densities are plotted versus the galaxy luminosity at 3.6 $\mu$m. The DM halos of our lenticular galaxies follow the same relations of SPARC galaxies within the errors. Thus, these ETGs live in similar DM halos as LTGs, indicating that they may be fading spirals. This also suggests that the evolution from spirals to lenticulars do not necessarily involve major structural changes in the DM halo, as one may expect after a major merger. Indeed, we can exclude major mergers in the recent past of our galaxies (1–2 Gyr). The orbital times in their outermost regions range from 0.3 to 0.8 Gyr: since the velocity fields are reasonably symmetric, the gas disk must have remained relatively unperturbed for several orbital times.

7. Conclusion

We analyzed the dynamics of three S0s that have inner CO emission and outer, extended HI disks. This special feature allowed us to trace the rotation curve from the innermost parts of the galaxies out to 10–20 effective radii. We derived mass models using Spitzer photometry at 3.6 $\mu$m and performed DM halo fits. Then, we compared the properties of our ETGs with those of LTGs from the SPARC sample to constrain the evolutionary history of lenticular galaxies. We found the following results:

1. The rotation-curve shapes of our ETGs are similar to those of LTGs of similar mass and surface brightness. This is consistent with a smooth transition from spirals to lenticulars.
2. The dynamically inferred values of the stellar mass-to-light ratio are relatively small for passive ETGs. The retrieved values are similar to those of LTGs (0.5–0.7 $M_\odot/L_\odot$). This suggests recent star formation events in these galaxies, consistent with the picture of spirals slowly fading into lenticulars.
3. The DM halo parameters follow the same scaling relations with galaxy luminosity as those of LTGs. This suggests that the transition from LTGs to ETGs happened without altering the halo structure of the galaxies, for example, via a violent event like a major merger.
4. The rotation curve of NGC 3626 has a significant inner decline driven by the stars and a subsequent rise driven by the halo, which is poorly fit by cuspy DM profiles. This suggests that DM cores may exist in high-mass galaxies.
5. We confirm that ETGs follow the same radial acceleration relation as LTGs, pointing to the universality of this relation for galaxies of different morphologies and evolutionary stages.

Our results paint a consistent picture of the evolution of lenticular galaxies, which is in agreement with previous works. A bigger sample of gas-bearing ETGs is needed to make more general statements. Such a sample might be accessible in the near future thanks to new CO surveys with ALMA and NOEMA, as well as HI surveys with the SKA and its pathfinders.

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Appendix A: Notes on individual galaxies

In this appendix, we present in detail the results of our 3D kinematic modeling with GALMOD and 3D BAROLO (see Sect. 3). For each galaxy, we describe intensity maps, position-velocity (PV) diagrams, and channel maps for both HI and CO emissions. The figures are structured in a uniform way.

A.1. NGC 2824

The observed HI intensity map is well reproduced by our model (Fig. A.1, top panels) apart from an over-density in the northern half of the disk, visible in the residual map, that cannot be reproduced by an axisymmetric disk. The observed PV diagrams display negative signal toward the galaxy center, associated to HI absorption, that our model cannot reproduce (Fig. A.1, bottom panels). Apart from these shortcomings, the HI kinematics are well reproduced. This is confirmed by Fig. A.2 comparing HI channel maps from the observed and model cubes.

The observed CO density maps indicates that the molecular disk is lopsided, extending farther toward the approaching part of the galaxy. This feature cannot be fully reproduced by our axisymmetric model (Fig. A.3, top panels). This also causes asymmetries in the PV diagrams (Fig. A.3, bottom panels). In our modeling, we give more weight to the extended, approaching half of the CO disk. Figure A.4 shows that the observed CO channel maps are well reproduced by our model.

A.2. NGC 3626

NGC 3626 is connected to the companion galaxy UGC 6341 by a diffuse HI bridge (Fig. A.5, top panels). We model only the HI emission associated to the disk of NGC 3626. The residual map shows that our axisymmetric model reproduces the HI disk very well, except for a ring feature at $R \approx 3'$. The major-axis PV diagram shows a prominent dip (Fig. A.5, bottom panels): since this feature is observed on both sides of the galaxy, we consider it an intrinsic feature of the rotation curve rather than an effect of noncircular motions. Overall, the HI kinematics is well reproduced, as shown by the channel maps in Fig. A.6.

The CO distribution is quite asymmetric (Fig. A.7, top panels), but this is likely due to the fact that the CO observations did not cover the full frequency range of galaxy emission, as can be appreciated by the sharp cutoff of the PV diagrams at low velocities (Fig. A.7, bottom panels). In our modeling, we manually adjust the rotation curve to reproduce the receding part of the galaxy. The resulting model also reproduces the approaching half, including the nonmeasured part of it. The channels maps in Fig. A.8 further demonstrate that the CO kinematics is well reproduced by our model.

The CO data confirms the strong decline in the inner parts of the rotation curve. Even in the extreme case of an edge-on CO disk (which is clearly not the case as seen in Fig. A.7), the CO rotation velocities at small radii are higher than the HI rotation velocities at large radii. Our rotation curve is about 50 km s$^{-1}$ higher in the inner parts than the one presented in Mazzei et al. (2014). We stress the HI disk is warped and the variations of $i$ and PA with radius are not well constrained by the data. At small radii, we adopt the values of $i$ and PA suggested by the CO disk, then we assume a simple linear rise to the higher values of $i$ and PA suggested by the HI disk at large radii. Clearly, a different modeling of the warp may slightly change the shape of the rotation curve and the resulting stellar mass-to-light ratios.

A.3. UGC 6176

The HI distribution and kinematics are regular and symmetric, so they are well described by our axisymmetric model (Figs. A.9 and A.10). We note, however, that the HI disk is poorly resolved and beam-smearing effects are severe, leading to a very broad PV diagram along the minor axis.

The CO disk is even less resolved than the HI emission, so we have only one ring in our model (Fig. A.11). The beam is elongated along the east-west direction leading to nontrivial beam-smearing effects. We could not properly reproduce the axial ratio of the observed CO map for any value of the disk inclination. We suspect that the CO beam is not Gaussian, so our model cannot reproduce such extreme beam-smearing effects. In the end, we adopted a relatively high inclination of 85°, which gives the closest match to the observed CO map. At any rate, the observed PV diagrams are well described by our model. The CO channel maps are also reasonably reproduced (Fig. A.12).

Although the ATLAS$^{3D}$ collaboration denotes UGC 6176 as a warped galaxy, we find no strong evidence for it. The PA of the HI and the CO disks are consistent with each other and show no clear variation with radius, while the inclination of the CO disk cannot be properly assessed from the available data.
Fig. A.1. **Top panels**: HI intensity maps from the observed cube (left), model cube (middle) and residual cube (right) of NGC 2824. The cross and the dashed line indicate the galaxy center and position angle, respectively. The beam is shown in the bottom left corner. The physical scale is illustrated in the bottom right corner. The contours correspond to $3\sigma_{\text{pseudo}}$, $6\sigma_{\text{pseudo}}$, and so on. The outermost contour of the model is plotted in blue over the observed intensity map for comparison.

**Bottom panels**: position-velocity diagrams taken along the major and minor axes from the observed cube (left), model cube (middle) and residual cube (right). The white squares show our derived rotation curve. The outermost contour of the model is plotted in blue over the observed PV diagrams for comparison. The density contours in the PV diagrams correspond to $-1.5\sigma$, $-3\sigma$, $1.5\sigma$, $3\sigma$, $6\sigma$, and so on (solid), where $\sigma$ is the measured noise in the cube (see Table 2).
Fig. A.2. HI channel maps for NGC 2824 from the observed cube (gray scale) and the model cube (blue contours). The line-of-sight velocity is shown in the upper left corner. The cross indicates the galaxy center. The contours correspond to $-\sigma$, $-1.5\sigma$ (dashed) and $1.5\sigma$, $3\sigma$, $6\sigma$, and so on (solid). In the first panel, the red contours correspond to arbitrary isophotes of the Spitzer image at 3.6 µm, the beam is shown in the bottom-left corner, and the physical scale is given in the bottom-right corner.
Fig. A.3. Same as Fig. A.1, but for the CO data of NGC 2824.
Fig. A.4. Same as Fig. A.2, but for the CO data of NGC 2824.
Fig. A.5. Same as Fig. A.1, but for the HI data of NGC 3626.
Fig. A.6. Same as Fig. A.2, but for the HI data of NGC 3626.
Fig. A.7. Same as Fig. A.1, but for the CO data of NGC 3626.
Fig. A.8. Same as Fig. A.2, but for the data CO of NGC 3626.
Fig. A.9. Same as Fig. A.1, but for the HI data of UGC 6176.
Fig. A.10. Same as Fig. A.2, but for the HI data of UGC 6176.
Fig. A.11. Same as Fig. A.1, but for the CO data of UGC 6176.
Fig. A.12. Same as Fig. A.2, but for the CO data of UGC 6176.
Appendix B: Posterior distributions of DM halo fits

In this appendix, we show the posterior distributions of the free parameters from DM halo fits with and without ΛCDM priors (see Sect. 5.1). When imposing ΛCDM priors (left panels), the posterior distributions are well-behaved and reasonably symmetric. Without imposing ΛCDM priors (right panels), the posterior distributions of the halo parameters become very complex, so their maximum-probability values are not fully trustworthy.

Fig. B.1. Posterior distributions for the DM halo fits of NGC 2824, with (left column) and without (right column) ΛCDM priors (see Sect. 5.1 for details). (a) NFW, (b) NFW (no priors), (c) Einasto, (d) Einasto (no priors), (e) DC14, and (f) DC14 (no priors).
Fig. B.2. Same as Fig. B.1, but for NGC 3626.
Fig. B.3. Same as Fig. B.1, but for UGC 6176.