Spectroscopic Observations of Obscured Populations in the Inner Galaxy: 2MASS-GC02, Terzan 4, and the 200 km s\(^{-1}\) stellar peak\(^*\)

Andrea Kunder\(^1\), Riley E. Crabb\(^1\), Victor P. Debbattista\(^2\), Andreas J. Koch-Hansen\(^3\), and Brianna M. Huhmann\(^1\)

\(^1\)Saint Martin’s University, 5000 Abbey Way SE, Lacey, WA 98503, USA
\(^2\)Jeremiah Horrocks Institute, University of Central Lancashire, Preston, PR1 2HE, UK
\(^3\)Zentrum für Astronomie der Universität Heidelberg, Astronomisches Rechen-Institut, Mönchhofstr. 12-14, D-69120 Heidelberg, Germany

Received 2021 February 25; revised 2021 June 2; accepted 2021 June 4; published 2021 August 3

Abstract

The interpretation of potentially new and already known stellar structures located at low latitudes is hindered by the presence of dense gas and dust, as observations toward these sight lines are limited. We have identified Apache Point Observatory Galaxy Evolution Experiment (APOGEE) stars belonging to the low-latitude globular clusters 2MASS-GC02 and Terzan 4, presenting the first chemical element abundances of stars residing in these poorly studied clusters. As expected, the signature of multiple populations coexisting in these metal-rich clusters is evident. We rederive the radial velocity of 2MASS-GC02 to be \(-87 \pm 7\) km s\(^{-1}\), finding that this cluster’s heliocentric radial velocity is offset by more than 150 km s\(^{-1}\) from the literature value. We investigate a potentially new low-latitude stellar structure and a kiloparsec-scale nuclear disk (or ring) that has been put forward to explain a high-velocity \((V_{\text{GSR}} \sim 200\) km s\(^{-1}\)) peak reported in several Galactic bulge fields based on the APOGEE commissioning observations. New radial velocities of field stars at \((l, b) = (-6^\circ, 0^\circ)\) are presented and combined with the APOGEE observations at negative longitudes to carry out this search. Unfortunately no prominent \(-200\) km s\(^{-1}\) peak at negative longitudes along the plane of the Milky Way are apparent, as predicted for the signature of a nuclear feature. The distances and Gaia EDR3 proper motions of the high-\(V_{\text{GSR}}\) stars do not support the current models of stars on bar-supporting orbits as an explanation of the \(+200\) km s\(^{-1}\) peak.

**Unified Astronomy Thesaurus concepts:** Galactic bulge (2041); Globular star clusters (656); Galaxy kinematics (602); Proper motions (1295); Milky Way formation (1053)

1. Introduction

Mapping the structure of the Galaxy is an essential endeavor, as the Milky Way’s (MW’s) assembly and star formation history is encoded in the kinematics, metallicities, ages, and spatial distribution of its stars (e.g., Freeman & Bland-Hawthorn 2002; Cooper et al. 2010; Minchev 2016; Helmi et al. 2018). The kinematics and composition of stars in the MW allow for the interpretation of individual components within the Galaxy and relate these to the fundamental properties of galaxy formation models.

The Galactic plane is a region that has remained particularly difficult to probe, as extreme obscuration by interstellar dust and gas, high source density, and confusion with foreground disk populations raise significant limitations in any kind of study (Gonzalez et al. 2012; Nataf et al. 2016; Alonso-García et al. 2018; Noguera-Lara et al. 2019). Consequently, our view of the inner Galaxy and Galactic bulge is dominated by observations of stars at Galactic latitudes larger than \(|b| \sim 2^\circ\).

Modern large surveys using new instrumentation on large telescopes are beginning to penetrate stellar populations that exist also along the plane of the bulge. For example, the infrared photometric VISTA Variables in the Vía Láctea (VVV) ESO Public Survey (Minniti et al. 2010) and the spectroscopic Apache Point Observatory Galaxy Evolution Experiment (APOGEE; Eisenstein et al. 2011) cover large areas of the plane (\(|b| < 2^\circ\)) of the bulge. This has allowed detailed studies of stars in this populous and crowded region in a more focused fashion.

However, one well-known gap in the literature concerns the population of globular clusters (GCs) in the bulge near the plane of the Galaxy. Surveys such as APOGEE have concentrated more on the GCs in the outskirts of the Galaxy than the bulge (e.g., Mészáros et al. 2020), with the bulge GCs needing to be treated in a more careful manner (e.g., Schiavon et al. 2017; Fernández-Trincado et al. 2019, 2020). This is unfortunate as the bulge GCs are important probes of the formation processes of the central parts of the Galaxy (e.g., Barbey et al. 2018). There is also a growing debate on the number of globular clusters that exist in the inner Galaxy, as discriminating real from spurious clusters candidates is troublesome (e.g., Froebrich et al. 2007; Gran et al. 2019; Rich et al. 2020).

Another ongoing issue concerns the so-called 200 km s\(^{-1}\) peak, which has been reported in the APOGEE fields closest to the plane of the Galaxy. This cold stellar feature \((\sigma_V \sim 35\) km s\(^{-1}\)) with a galactocentric velocity (\(V_{\text{GSR}}\)) of \(~200\) km s\(^{-1}\) was first reported in several Galactic bulge fields based on the APOGEE commissioning observations (Nidever et al. 2012). Nidever et al. (2012) originally proposed that the cold high-\(V_{\text{GSR}}\) stars are on bar orbits.

Both Li et al. (2014) and Gómez et al. (2016) showed that bar-supporting orbits would produce a high-velocity shoulder, but no discrete peak. Although indeed, many of the APOGEE fields exhibit shoulders and not peaks, 3 of the 53 of the APOGEE fields at positive longitudes do show a discrete \(V_{\text{GSR}}\) peak with a clear trough (Zhou et al. 2017). For the negative longitudes, Zhou et al. (2021) find no distinct cold high-\(V_{\text{GSR}}\) peaks in the APOGEE fields, instead finding that complex velocity distributions fit these fields best. However,
observations at the negative longitudes are not as extensive or complete, hampering a detailed study.

Aumer & Schönrich (2015) argued that the APOGEE selection function favors young stars, and younger stars would be trapped by the bar into resonant orbits and could give rise to the cool, high-velocity peaks. However, Zasowski et al. (2017) showed that stars in the APOGEE high-$V_{\mathrm{GSR}}$ peaks do not exhibit distinct chemical abundances or ages, indicating that they are not predominantly comprised of younger stars. Recently, McGough et al. (2020) showed that stars in the propeller orbit family, in isolation, have a kinematic signature similar to that of the 200 km s$^{-1}$ peak.

A further explanation of the cold, high-velocity peak is that the stars in the high-$V_{\mathrm{GSR}}$ peak are the signature of a kiloparsec-scale nuclear disk (or ring) that could exist at the center of the Milky Way (Debattista et al. 2015). In this interpretation, the high-velocity peak seen clearly in the APOGEE data at $(l, b) = (6^\circ, 0^\circ)$ is the tangent point of a nuclear structure supported by $x_2$ orbits, which are perpendicular to the bar (see, e.g., Figure 1 in Debattista et al. 2015). We stress that the kiloparsec-scale nuclear disk put forward by this model is unrelated to the nuclear stellar disk that could exist at the center of the Milky Way (Debattista et al. 2015).

The results presented in this paper come from the APOGEE giant stars contained within the grid shown in Figure 1. These observations cover primarily the plane of the inner bulge ($|b| < 1^\circ$). The grid indicates our divisions used in Section 4 (in particular the $V_{\mathrm{GSR}}$ histograms shown in Figure 6) to search for potential high-$V_{\mathrm{GSR}}$ peaks.

2.2. Anglo-Australian Telescope Observations

One bulge field at $(l, b) = (-6^\circ, 0^\circ)$ was observed on 2020 June 21 (NOAO PropID: 2020A-0368; PI: A. Kunder) using the AAOmega multifiber spectrograph on the Anglo-Australian Telescope (AAT). As this spectrograph is a dual-beam system,
The BEAM extinction calculator will be located in the software Team 2015 giants (conditions, our exposure time was 2 × 30 minutes, after which the clouds rolled in. Only the spectra from the Red arm were used, as this is where the prominent Calcium Triplet lines could still be distinguished. These spectra covered the wavelength used, as this is where the prominent Calcium Triplet lines could still be distinguished. These spectra covered the wavelength range from ~0.1–0.7 mag in both ΔE(J−K) and ΔAK.

The AAT spectra were cross correlated using the IRAF cross-correlation routine, xcsaoa, using the calibration stars taken with the identical setup. Our calibration stars include the following APOGEE stars: 2M18264551-1747096 (HRV = −69.2 km s⁻¹), 2M18255968-174935 (HRV = 49.6 km s⁻¹), 2M18205442-1751063 (HRV = 97.4 km s⁻¹), 2M18235228-1653096 (HRV = −52.7 km s⁻¹), and 2M18233429-1841586 (HRV = 165.8 km s⁻¹), which were observed in a previous observing run (see Kunder et al. 2020a), as well as the APOGEE stars that are located in the same field as our newly observed stars: 2M17342312-3335162 (HRV = −2.2 km s⁻¹) and 2M17340454-3332161 (HRV = −67.3 km s⁻¹). In general, the consistency of our velocity results are 8 km s⁻¹, which is in agreement with the errors reported by xcsao for each individual measurement.

3. Globular Clusters

With the Gaia satellite, new astrometric data have become available for approximately a billion stars in the MW galaxy (Gaia Collaboration et al. 2018a, 2021; Lindegren et al. 2020). Combining radial velocities with proper motions from Gaia allows for all three velocity components of stellar motion to be determined. The possibility of using the Gaia proper motions, therefore, completes the measurements of the space motions of a star.

Gaia is ongoing with an anticipated five-year mission lifetime, and the most recent data release (EDR3; Gaia Collaboration et al. 2021) is based on 34 months of observations. Especially in crowded regions of the sky, Gaia faces challenges with astrometric solutions (e.g., Sanders et al. 2019); this will undoubtedly improve by end of mission. Here we corroborate the accuracy of Gaia proper motions in the plane of the Galaxy near the bulge by checking the consistency of the proper motions of stars in GCs. Six globular clusters in the APOGEE footprint located in or close to the plane were used for this purpose; five explored here and one more taken from Kunder & Butler (2020b).

Table 1 lists the individual globular clusters used to verify the reliability of the Gaia EDR3 proper motions. Column 1 is the name of the GC, column 2 lists the average proper motion for the stars selected as cluster candidates, column 3 lists heliocentric distances as reported in the 2010 edition of Harris (1996), column 4 lists the radial velocity of the cluster, and columns 5 and 6 are the published proper motions for the GCs as reported by Rossi et al. (2015) and Vasiliev & Baumgardt (2021), respectively.

| Cluster Name | Average Proper Motion (mas yr⁻¹) | Distance (pc) | Radial Velocity (km s⁻¹) | Published Proper Motion (mas yr⁻¹) |
|--------------|----------------------------------|---------------|--------------------------|----------------------------------|
| Terzan 1     | 0.3 ± 0.1                       | 12.5 ± 0.5    | −74.6 ± 2.6              | −73.0 ± 2.0                      |
| Terzan 2     | 0.4 ± 0.2                       | 11.3 ± 0.7    | −69.2 ± 2.8              | −68.0 ± 2.5                      |
| Terzan 3     | 0.2 ± 0.1                       | 13.0 ± 0.8    | −52.7 ± 3.0              | −51.0 ± 3.0                      |
| Terza 4      | 0.1 ± 0.0                       | 12.8 ± 0.6    | −49.6 ± 1.8              | −48.0 ± 1.5                      |
| Terzan 5     | 0.3 ± 0.2                       | 13.2 ± 0.9    | −97.4 ± 3.2              | −96.0 ± 3.0                      |
| Terzan 6     | 0.2 ± 0.1                       | 14.0 ± 0.7    | −69.2 ± 2.6              | −68.0 ± 2.5                      |

Previous papers to isolate globular cluster stars use proper motion as a criterion for membership (e.g., Schiavon et al. 2017; Baumgardt et al. 2019; Vasiliev 2019a; Horta et al. 2020), so we searched for cluster stars in Palomar 6, Terzan 2, Terzan 4, Liller 1, and 2MASS-GC02 using primarily APOGEE radial velocities and the radial distance of a star from the center of the cluster. Also, because globular cluster stars are in general more metal poor than the field population (e.g., Harris 1996; Schiavon et al. 2017; Horta et al. 2020; Mészáros et al. 2020), the metallicities of the stars were used to select candidate cluster members. The published metallicity and radial velocity values for the clusters were used as soft boundaries to guide selection.

Twenty three cluster stars across the five GCs were identified based on their radial velocity, distance from the center of

![Figure 2. The dereddened 2MASS color–magnitude diagram showing the 130 giants (triangles) for which radial velocities have been determined from AAOmega@AAT. The APOGEE stars observed at (l, b) = (−5°5, 0°) are also shown (gray bold points), as are the underlying 2MASS stellar distribution in this field (small gray points). All stars have been dereddened using the extinctions from Gonzalez et al. (2012) and the Nishiyama et al. (2009) extinction law. The uncertainties in the reddening and extinction law dominate the errors in this plot and range from ~0.1–0.7 mag in both ΔE(J−K) and ΔAK.](image)
cluster, and, when available, metallicity. Table 2 lists the individual stars in each targeted globular cluster, as well as the radial velocity and the elemental abundances [Fe/H], [C/Fe], [N/Fe], [Na/Fe], [Mg/Fe], and [Al/Fe] of the star as reported by APOGEE DR16. The distances in column 4 are from StarHorse, taken directly from Queiroz et al. (2020). While APOGEE stars in the cluster Palomar 6 were studied previously in Schiavon et al. (2017) and Mészáros et al. (2020), we still carried out a selection of stars in this cluster. This was to avoid any kind of potential proper motion biases that would artificially lower the scatter in the average proper motions of these stars. We present the first APOGEE stars in the globular clusters 2MASS-GC02, Terzan 4 and Liller 1.

### 3.1. 2MASS-GC02

Figure 3 shows an example of the selection of APOGEE stars in an understudied GC, 2MASS-GC02 (Hurt et al. 2000). An overdensity of stars close to the center of the cluster with radial velocities of ~87 km s\(^{-1}\) and metallicities of [Fe/H] ~ −1.0 is apparent. These are likely not field stars, but instead almost certainly belong to 2MASS-GC02. Curiously Borisssova et al. (2007) present 15 stars with S/N > 30 that lie within 1/3 of 2MASS-GC02 and that have colors and magnitudes consistent with the giant branch of 2MASS-GC02. From 12 of these stars, they found an average radial velocity of ~−238 ± 36 km s\(^{-1}\). However, Figure 3 shows no stars with such large radial velocities.

The right panel of Figure 3 shows that our seven candidate 2MASS-GC02 stars also have colors and magnitudes consistent with being members of the cluster. The reddening of this cluster is extreme, A\(_p\) > 15 mag or \(E(J-K) = 3.1 ± 0.5\) mag (Hurt et al. 2000; Ivanov et al. 2000), and it has been noted that the extinction toward this cluster is highly variable (up to 50% change in extinction over small regions) and with significant deviations from the standard extinction law (Alonso-García...
et al. 2015). Peñaloza et al. (2015) have found the stars in this cluster are $\alpha$ enhanced, as is typical of old globular clusters in the bulge (e.g., Dias et al. 2016; Mészáros et al. 2020; Horta et al. 2020). The BaSTI (Pietrinferni et al. 2004, 2006) $\alpha$-enhanced stellar evolution models5 were adopted to indicate the approximate location of the red giant branch of 2MASS-GC02. We used the publicly available BaSTI isochrone that best matches the cluster’s observed parameters (see text). The APOGEE uncertainties are in $[\text{Fe}/\text{H}]$ are 0.1–0.2 dex, the APOGEE uncertainties in radial velocity are $<0.5$ km s$^{-1}$, and the uncertainties in the dereddened CMD range from $\sim0.1$–0.7 mag in both $\Delta E(J − K)$ and $\Delta A_K$.

Values dramatically differ in the zero point by 3–11 mas yr$^{-1}$, and no correlation between the Gaia EDR3 proper motions and those presented by Rossi et al. (2015) are apparent. Although systematic errors in Gaia proper motions have been reported (e.g., Vasiliev 2019b; Lindegren et al. 2018; Gaia Collaboration et al. 2018b), offsets of $>1$ mas yr$^{-1}$ are not thought to be likely.

3.2. Chemical Element Abundances

Observational evidence demonstrates that most Galactic GCs host (at least) two main groups of stars with different chemical composition. One population of stars, referred to as the first generation, have a similar chemical composition as halo field stars with similar metallicity, while the other population of stars have enhanced helium, nitrogen, and sodium abundances, but are depleted in carbon and oxygen (e.g., Kraft 1994; Carretta et al. 2009; Marino et al. 2019). The origin of these multiple populations is still an open issue (e.g., Gratton et al. 2012; Renzini et al. 2015; Bastian & Lardo 2018).

GCs in the inner Galaxy represent a moderately metal-rich population and are thought to be some of the oldest GCs in the Galaxy (e.g., VandenBerg et al. 2013; Massari et al. 2019). However, bulge GCs are less studied than those in the halo, as high and often variable extinction, as well as crowding and the difficulty in separating true bulge stars from intervening thin and thick disk stars, makes them more difficult to analyze in detail (e.g., Bica et al. 2016). Whereas almost all GCs studied have shown signs of multiple populations with certain dependencies on mass (and age), it is less clear how $[\text{Fe}/\text{H}]$ metallicity affects the multiple-population phenomena (Kayser et al. 2008; Carretta et al. 2009; Hanke et al. 2017; Bastian & Lardo 2018). For example, old, metal-rich clusters have been found that do not exhibit multiple populations (e.g., Salinas & Strader 2015) and these clusters have been used to suggest improvement in the models of GC formation (Tang et al. 2017).

This paper is the first to present the APOGEE stars and their chemical element abundances in the clusters 2MASS-GC02 and Terzan 4. As such, a search for multiple populations in these clusters is carried out. Due to the small numbers of stars per cluster with elemental abundances, we are hesitant to carry out a deeper analysis, e.g., identifying gaps in the elemental distribution or identifying clusters that are chemically

---

5 http://basti-iac.oa-abruzzo.inaf.it
of stars within $\sim 20^\circ$ of the Galactic Center. Many of these peaks were shown to not be statistically significant (Zhou et al. 2017) and there has also been the suggestion that APOGEE’s selection function was biased (Aumer & Schönrich 2015). Other studies have shown that some of these velocity peaks or shoulders can be associated with resonant bar orbits (Molloy et al. 2015; McGough et al. 2020). It has also been suggested that the most prominent velocity peak at $(l,b) = (6^\circ,0^\circ)$ is caused by stars moving on $x_2$ orbits in a kiloparsec-sized nuclear disk or ring (Debattista et al. 2015, 2018).

4.1. Velocity Distributions

When Nidever et al. (2012) first reported the so-called 200 km s$^{-1}$ peak at positive longitudes along the plane of the Galaxy, there were no observations at the negative longitudes to search for a counterfeature (see Figure 1). This was unfortunate, because its prominence (or “peakiness”) and how widespread it is at negative longitudes, could help discriminate between scenarios which attempt to explain this feature. For example, if the counterfeature is widespread and appears as a shoulder instead of a discrete peak, this would agree with the models put forth by e.g., Li et al. (2014), Aumer & Schönrich (2015), Gómez et al. (2016), and Zhou et al. (2021) that the 200 km s$^{-1}$ peak is composed of stars on bar-supporting orbits. If the counterfeature is confined to longitudes between $1 \sim 6^\circ$ and $1 \sim 8^\circ$ and is a discrete peak, this would agree with the suggestion of Debattista et al. (2018) that the 200 km s$^{-1}$ peak is composed of stars belonging to a nuclear disk or ring.

Figure 6 shows the galactocentric velocities for the APOGEE DR16 at negative longitudes compared to those from APOGEE DR12/DR14 for the positive longitudes. Our new observations described in Section 2 are also folded in, as they provide a more direct comparison to the $(l,b) = (6^\circ,0^\circ)$ field than using the APOGEE DR16 observations alone, as there are no APOGEE stars at longitudes between $+6^\circ$ and $+7^\circ$.

A nuclear feature would be present mainly within $b \pm 1^\circ$ and also mainly around $|l| \sim 5^\circ-8^\circ$, as this is the line-of-sight tangent to the nuclear disk/ring (e.g., see Figure 1; Debattista et al. 2015). The left panel of Figure 6 concentrates on this region of the bulge. The right panel shows the velocity distributions for stars at longitudes reaching to $|l| = 10^\circ$ and $|l| = 2^\circ.5$ as well as fields $\pm 2^\circ$ from the plane of the Galaxy. Any high-$V_{\mathrm{GSR}}$ features/signatures in these fields would not be due to a nuclear feature in these off-plane fields.

A potential counterpeak at $\sim 200$ km s$^{-1}$ appears to become more prominent and more separated from the rest of the distribution as the observations approach $|l| = 6^\circ$, but the $1 \sim +6^\circ$ observations show a significantly more pronounced peak as compared to at $1 \sim 6^\circ$. Our radial velocity distributions show that stars do occupy the high-$V_{\mathrm{GSR}}$ regime also around $(l,b) = (5^\circ.5,0^\circ)$, but there is no obvious negative high-$V_{\mathrm{GSR}}$ peak.

Figure 7 shows the APOGEE DR16 combined with the AAT velocities for 312 stars in the field centered on $(l,b) = (5^\circ.5,0^\circ)$. A Shapiro–Wilk test for normality on the radial velocity distribution gives $p$-values <0.05, indicating that the radial velocity distribution formally deviates from a Gaussian distribution. A $k$-means clustering algorithm is used to determine where the radial velocity distribution is aggregated. A 4 peak curve provides the best fit to the $(l,b) = (5^\circ.5,0^\circ)$.
velocity distribution and a Gaussian mixture model (GMM) indicates that the radial velocity peaks are at $[-45 \text{ km s}^{-1}]$,
peaks are located for the fields shown in Figure 7. Although formally these statistical tests provide a quantitative characterization of the velocity distribution of stars near the plane of the Galaxy, we believe the large velocity dispersion combined with the sample size makes it difficult to conclusively determine the underlying velocity distribution.

It is evident that any preference for an excess of stars at \( V_{\text{GSR}} \sim -200 \text{ km s}^{-1} \) in the \( l = -5^{\circ}5 \) field is not as extreme as was seen in the \( l = +5^{\circ}5 \) field, nor is it confined to only this field. For example, the field at \((l,b)=(−10^{\circ}, 0^{\circ})\) at \((l,b)=(−4^{\circ}, 0^{\circ})\) also may show some indication for a \( V_{\text{GSR}} \sim -200 \text{ km s}^{-1} \) peak.

The Figure 6 presented here is similar to the velocity distributions shown in (Zhou et al. 2021, their Figure 2), except that we have \( \sim100 \) more stars with radial velocities centered on \((l,b)=(−6^{\circ},0^{\circ})\) due to our new AAT observations. We also show velocity histograms spanning ranges of longitude to specifically explore a counterfeature at negative longitudes. Instead of selecting 2° fields based on where the positive longitude fields are centered, we separate the stars so they naturally follow the sparser coverage along the negative longitudes (see Figure 1). In particular, as there is no negative longitude field centered on \((l,b)=(−6^{\circ},0^{\circ})\), we do not show a velocity distribution for this field. Instead, the \((l,b)=(−5^{\circ}8,0^{\circ})\) field presented here spans half a degree from \(-5^{\circ}5\) to \(-6^{\circ}0\) and the \((l,b)=(−5^{\circ}3,0^{\circ})\) field spans half a degree from \(-5^{\circ}0\) to \(-5^{\circ}5\). These fields can be directly compared to the APOGEE stars at positive longitudes and hence more closely mirrors the APOGEE observations in hand.

### 4.2. Velocity and Distance Correlations

Li et al. (2014) and Zhou et al. (2017) use MW simulations to remark that in a distance-velocity diagram, higher-velocity particles should be at larger distances. This is due to the geometric intersections between the line of sight and the particles orbit; at the tangent points of a star’s orbit, it will have the highest velocity and remain there longer relative to us. The StarHorse distances (Queiroz et al. 2020) released for the APOGEE DR16 stars can be used to explore if the high-velocity stars seen in the histograms presented above are consistent with particle motions in an axisymmetric disk.

Figure 8 shows the velocity distributions for those stars with StarHorse distances between 4 and 6 kpc compared to the velocity distributions for those stars with distances >6 kpc. In both fields there are still stars with distances between 4 and 6 kpc that occupy the high-\( V_{\text{GSR}} \) peaks. There is only a mild trend for stars with on average larger distances to have larger velocities. The stars with distances between 4 and 6 kpc have dominant velocity peaks at \( \sim50 \text{ km s}^{-1} \). In contrast, stars with distances larger than 6 kpc are not as strongly peaked and instead have tails extending to larger velocities. This suggests that the high-\( V_{\text{GSR}} \) peak cannot be explained by stars at larger distances alone, although it may also be that the StarHorse distances are not reliable for these stars. The plane of the Galaxy has especially high and variable extinction values and this may affect the accuracy of the StarHorse distances.

### 4.3. Velocity and Proper-motion Correlations

Gaia proper motions are available for a number of stars along the plane and in Section 3 we showed that they can be used also in this crowded region of the Galaxy. In particular, from an
Table 3
Statistical Properties of the APOGEE Velocity Distributions

| Field         | $l$ Range | $b$ Range | Num Stars | # Gaussians | Mean      | Sigma     | Weight     | Data         |
|---------------|-----------|-----------|-----------|-------------|-----------|-----------|------------|--------------|
| ($l, b$) = (-10, 0) | -11:-9   | -1:1      | 257       | 2           | -71, -177 | 59, 67    | 0.70, 0.30  | APOGEE DR16  |
| ($l, b$) = (-5.5, 0) | -6:-5    | -1:1      | 312       | 4           | -45, -133, 53, -257 | 36, 31, 43, 31 | 0.34, 0.31, 0.21, 0.14 | APOGEE + AAT |
| ($l, b$) = (-5.5, 0) | -6:-5    | -1:1      | 220       | 4           | -31, -122, 57, -256 | 29, 34, 34, 30 | 0.29, 0.33, 0.22, 0.15 | APOGEE       |
| ($l, b$) = (-5.8, 0) | -6:-5.5  | -1:1      | 131       | 4           | -53, -140, 49, -260 | 37, 29, 54, 33 | 0.39, 0.32, 0.17, 0.12 | APOGEE + AAT |
| ($l, b$) = (-5.3, 0) | -5.5:-5  | -1:1      | 181       | 4           | -121, -30, 57, -255 | 34, 29, 35, 30 | 0.34, 0.26, 0.24, 0.16 | APOGEE DR16  |
| ($l, b$) = (-5.5, -2) | -6:-5    | -3:-1     | 116       | 2           | -25, -101 | 73, 162   | 0.70, 0.30  | APOGEE DR16  |
| ($l, b$) = (-5.5, 2) | -6:-5    | 1:3       | 86        | 3           | -87, 43, -254 | 43, 56, 39 | 0.55, 0.32, 0.13 | APOGEE DR16  |
| ($l, b$) = (-4.5, 0) | -5:-4     | -1:1      | 245       | 2           | -24, -209 | 76, 76    | 0.72, 0.28  | APOGEE DR16  |
| ($l, b$) = (-3.5, 0) | -4:-3     | -1:1      | 80        | 3           | -62, -222, 94 | 54, 62, 46 | 0.42, 0.37, 0.21 | APOGEE DR16  |
| ($l, b$) = (-2.0) | -3:-1     | -1:1      | 401       | 4           | -99, 4, 129, -228 | 47, 47, 74, 52 | 0.37, 0.33, 0.16, 0.15 | APOGEE DR16  |
| ($l, b$) = (+2.0) | 1:3       | -1:1      | 477       | 3           | 8.7, 148, -126 | 53, 71, 60 | 0.49, 0.30, 0.21 | APOGEE DR16  |
| ($l, b$) = (+3.5, 0) | 3:4       | -1:1      | 67        | 2           | 2, 144 | 62, 77    | 0.73, 0.27  | APOGEE DR12/14 |
| ($l, b$) = (+4.5, 0) | 4:5       | -1:1      | 187       | 2           | 32, 108 | 83, 160   | 0.72, 0.28  | APOGEE DR12/14 |
| ($l, b$) = (+5.3, 0) | 5:5.5     | -1:1      | 126       | 3           | 18, 144, -137 | 44, 80, 119 | 0.66, 0.29, 0.05 | APOGEE DR12/14 |
| ($l, b$) = (+5.8, 0) | 5:5.6     | -1:1      | 126       | 2           | 19, 233 | 71, 50    | 0.79, 0.21  | APOGEE DR12/14 |
| ($l, b$) = (+5.5, 0) | 5:6       | -1:1      | 252       | 2           | 24, 236 | 75, 43    | 0.86, 0.14  | APOGEE DR12/14 |
| ($l, b$) = (+5.5, -2) | 6:5       | -3:-1     | 899       | 2           | 16, 92 | 74, 112   | 0.63, 0.37  | APOGEE DR16   |
| ($l, b$) = (+5.5, 2) | 6:5       | 1:3       | 618       | 2           | 8.2, 117 | 82, 94    | 0.63, 0.37  | APOGEE DR16   |
| ($l, b$) = (+10, 0) | 9:11      | -1:1      | 253       | 2           | 50, 158 | 57, 125   | 0.75, 0.25  | APOGEE DR12/14 |
stars with larger distances

0.15 mas \ yr^{-1}

analysis of globular cluster stars, a star-to-star scatter of 0.15 mas \ yr^{-1} or less can be expected. Formally the precision of the EDR3 proper motions is \( \sim 0.07 \) mas \ yr^{-1} (at \( G = 17 \)) for Gaia EDR3 as compared to 0.2 mas \ yr^{-1} (at \( G = 17 \)) in Gaia DR2 (Lindegren et al. 2020). Gaia EDR3 proper motions offer not only an increase in precision, but the number of stars with proper motions along the plane also increases by \( \sim 15\% \) as compared to Gaia DR2. Therefore, our sample of stars along the plane is useful to probe the correlation between the high-\( V_{\text{GSR}} \) regime also in proper-motion space.

Figure 9 shows the correlation between velocity and \( \mu_\ell \) for 249 stars at \( (l,b) = (5^\circ,5.0^\circ) \). At a velocity of \( V_{\text{GSR}} > 100 \) km \ s\(^{-1} \), stars have proper motions with larger \( |\mu_\ell| \) values than the rest of the sample. In particular, stars with \( >100 \) km \ s\(^{-1} \) have \( \mu_\ell \) of \( \sim 7.5 \) mas \ yr^{-1} as compared to \( \mu_\ell \) of \( -4 \) mas \ yr^{-1} for the bulge field.

The \( \mu_\ell \) for 328 stars at \( (l,b) = (5^\circ,5.0^\circ) \) shows the same trend for the positive high-\( V_{\text{GSR}} \) stars—the stars with \( \sim 200 \) km \ s\(^{-1} \) velocities have \( |\mu_\ell| \) proper motions larger than the rest of the sample. In contrast, the negative high-\( V_{\text{GSR}} \) stars have a similar \( \mu_\ell \) to the bulge field. There is no preponderance for the negative high-\( V_{\text{GSR}} \) regime to have different \( \mu_\ell \) proper motions to the bulge field.

The same trend in proper motion is also seen \( \pm 2^\circ \) from the plane. There is no indication that the stars along the plane of the Galaxy at \( |l| = 5^\circ,5 \) differ in \( \mu_\ell \) as compared to those \( \pm 2^\circ \) from the plane at \( |l| = 5^\circ,5 \).

McGough et al. (2020) use 71 stars at \( l \sim 6^\circ \), 52 stars at \( l \sim 8^\circ \) and 27 stars at \( l \sim 4^\circ \) to show that the high-\( V_{\text{GSR}} \) stars occupy a smaller range in \( \mu_\ell \) than the full sample of stars. They attribute this to stars on periodic orbits in an exponential bar, such as propellor orbits. They were not able to investigate the proper motions at negative longitudes. Zhou et al. (2021) corroborate the result from McGough et al. (2020) result using DR2 proper motions for the APOGEE DR16 sample. Here we are able to offer additional insights on the correlation between velocity and \( \mu_\ell \) in the bulge using the better precision and sampling of the EDR3 proper motions combined with both the APOGEE and our new AAT observations.

Besides Gaia proper motions, proper motions of bulge stars from Smith et al. (2018) are publicly available through the VIRAC catalog, a near-infrared proper motion and parallax catalog of the VISTA Variables in the VVV survey. Out of the 3614 APOGEE stars within \( |b| < 1^\circ \), \( (J-K_s) > 0.5 \) and log \( g > 3.8 \), about half of those (1656) have a VIRAC proper motion and about a third of those (592) have a flag = 1, indicating it is reliable. In particular, in the \((-5,5,0)\) field, there are 220 stars with VIRAC proper motions and in the \((+5,5,0)\) field there are 98 stars with VIRAC proper motions. This overlap is smaller then when using Gaia EDR3. There are 1510 APOGEE stars within \( |b| < 1^\circ \), \( (J-K_s) > 0.5 \) and log \( g > 3.8 \) that have a Gaia EDR3 proper motion; \( 87\% \) of these have absolute proper motion uncertainties \( < 1 \) mas \ yr^{-1}. The number of APOGEE stars with Gaia proper motions in the \((-5,5,0)\) and \((5,5,0)\) field is greater by a factor of 1.5–2.

Using the VIRAC, proper motions result in the same trends as seen above in Figure 9 using the more extensive Gaia EDR3 proper motions. APOGEE stars with radial velocities of \( \sim 200 \) km \ s\(^{-1} \) have larger \( |\mu_\ell| \) values. This is the case both for the bulge stars with \( |b| < 1^\circ \) as well as the stars at \( |b| \sim 2^\circ \).

We prefer not to combine the Gaia and VIRAC proper motion catalogs due to the offsets between these catalogs caused by the drift motion of the pool of reference stars used for each pawprint set in VIRAC (see e.g., Smith et al. 2018; Clarke et al. 2019). These are estimated to be \( +2.20 \) mas \ yr^{-1} in \( \mu_\ell \) and \( -4.85 \) mas \ yr^{-1} in \( \mu_\ell \); the authors mention it is quite probable that there are unknown systematic uncertainties besides these offsets as well.

### 5. Discussion

Proper motions from Gaia EDR3, radial velocities from APOGEE, distances from StarHorse, and our own observations from the AAT are combined for stars along the plane of the Galaxy with the aim of testing the hypothesis that the
high-$V_{\text{GSR}}$ stars are the signature of a nuclear disk or ring (Debattista et al. 2018). Our new observations from the AAT at $(l,b)=(-6^\circ,0^\circ)$ are combined with the APOGEE DR16 observations to increase the sample size of stars with radial velocities at negative longitudes where the tangent point of a nuclear feature should be evident.

From a GMM analysis the putative negative high-$V_{\text{GSR}}$ peak at $(l,b)=(-5^\circ.5,0^\circ)$ is at $V_{\text{GSR}} \sim -250 \text{ km s}^{-1}$ whereas the positive high-$V_{\text{GSR}}$ peak at $(l,b)=(5^\circ.5,0^\circ)$ is at $V_{\text{GSR}} \sim 230 \text{ km s}^{-1}$. Therefore, the negative high-$V_{\text{GSR}}$ stars have the larger $|V_{\text{GSR}}|$, in contrast to what is expected for a nuclear feature.

Along the plane of the Galaxy, the new Gaia EDR3 proper motions increase the number of APOGEE stars with proper motions by $\sim 15\%$ as compared to Gaia DR2. The proper motions of the high-$V_{\text{GSR}}$ stars occupy a small range in proper motions, spanning $\sim 5 \text{ mas yr}^{-1}$. This is inconsistent with what is predicted for stars in a nuclear feature, as proper motions of stars in both a nuclear disk and nuclear ring would span a wider range of in $\mu_\ell$, $\sim 10 \text{ mas yr}^{-1}$ (Debattista et al. 2018).

The $\mu_\ell$ of the positive high $V_{\text{GSR}}$ exhibit larger $|\mu_\ell|$ proper motions than the bulge stars. However, the negative high-$V_{\text{GSR}}$ stars have similar $\mu_\ell$ proper motions as bulge stars. Zhou et al. (2021) argues that this observed $V_{\text{GSR}}-\mu_\ell$ distribution is consistent with the predictions of a simple MW bar model as presented in Shen et al. (2010). This is especially the case if the APOGEE stars were to lie predominantly in front of the Galactic Center. However, the Shen et al. (2010) model does not predict a distinct peak at $(l,b)=(5^\circ.5,0^\circ)$. The only way to produce this peak in the Shen et al. (2010) model is if the high-$V_{\text{GSR}}$ stars are at distances at $\sim 8.5 \pm 1 \text{ kpc}$ (see e.g., Figure 3 in Li et al. 2014). Accordingly there is a tension in the Shen et al. (2010) model between the APOGEE stars needing to be at $\sim 8.5 \text{ kpc}$ to explain any high-$V_{\text{GSR}}$ peak but to be primarily in front of the Galactic center to match the $\mu_\ell$ proper motions.

Although there are negative high-$V_{\text{GSR}}$ stars residing in the negative longitude fields, the negative high-$V_{\text{GSR}}$ peak is not nearly as prominent or discrete as seen in the positive high-$V_{\text{GSR}}$ peak at the positive longitudes (see Figure 6). Stars on propellor orbits as described by McGough et al. (2020) are not able to explain this. The propellor orbits can produce positive high-$V_{\text{GSR}}$ stars with similar proper motions to what is observed, but there is no reason why APOGEE would select just these orbits at $(l,b)=(5^\circ.5,0^\circ)$. Although it has been hypothesized that the APOGEE selection function preferably selects young stars (Aumer & Schönrich 2015), stars on propellor orbits would not necessarily be young and there has been no observational evidence to suggest that high-$V_{\text{GSR}}$ stars are in fact younger than the bulk of the population (Zasowski et al. 2016; Zhou et al. 2017, 2021). Therefore, those propellor orbits would give rise to a strong secondary peak only at $(l,b)=(5^\circ.5,0^\circ)$, which is not a satisfactory way of explaining a distinct high-$V_{\text{GSR}}$ peak.

It was shown in (e.g., Li et al. 2014) the high-$V_{\text{GSR}}$ stars can be explained by stars at larger distances, although this would produce a high-$V_{\text{GSR}}$ peak instead of a shoulder. Using StarHorse distances we are not able to confirm that the high-$V_{\text{GSR}}$ stars are at further distances. This leaves open a different interpretation for the peak in the APOGEE high-$V_{\text{GSR}}$ stars than current MW models of the orbits of bulge stars suggest.

Our analysis assumes that the completeness between APOGEE-2 and APOGEE for the bulge fields are similar without major differences in the selection effects. This is the case for other studies using APOGEE DR16 to compare the positive and negative longitudes of the inner galaxy (e.g., Zhou et al. 2021; Rojas-Arriagada et al. 2020). Zasowski et al. (2017) report that in both APOGEE-2 and APOGEE, bulge giants are selected based on their dereddened $(J-K)_0$ colors, requiring $(J-K)_0 > 0.5$, and Figure 2 shows no obvious indication that the APOGEE stars at negative longitudes do not probe the full $(J-K)_0$ color range for bulge giants. Although, there have

![Figure 9](image_url)
been no detailed investigations to investigate selection effects in APOGEE DR16 (see e.g., Nandakumar et al. 2017; Fragkoudi et al. 2018). Queiroz et al. (2021) uses cartesian density maps of the bulge stars to surmise that the complex mix of stellar populations and selection function of APOGEE does provide a homogeneous coverage of the entire inner Galaxy. Fortunately, the detection of a high-velocity peak at negative longitudes should be as apparent, if not more, using stars at negative longitudes of the Galaxy (Debattista et al. 2018). Yet still, our different cuts and stellar divisions within the APOGEE-2 data set (e.g., Figures 6 and 7; Table 3) do not show any obvious −200 km s⁻¹ peak.

6. Conclusions

The APOGEE DR16 observations combined with Gaia EDR3 proper motions are used to investigate the hypothesis that the cold high-velocity peak (at V_{GSR} ~ 200 km s⁻¹; Nidever et al. 2012) are due to a kiloparsec nuclear feature (Debattista et al. 2015, 2018). The APOGEE DR16 data in the negative longitudes were examined for a counterpart at V_{GSR} ~ −200 km s⁻¹. Some evidence for a high-velocity shoulder in a number of different APOGEE negative longitude fields along the plane is seen, but no dominant V_{GSR} ~ −200 km s⁻¹ feature is apparent. Unfortunately the sample size is small compared to the large velocity dispersion of the bulge population along the plane of the Galaxy and a larger sample size as well as wider coverage (e.g., l = −6° to −8°) would offer a deeper view on characterizing any ~ [200] km s⁻¹ feature. Distances and proper motions of the high-velocity stars are inconsistent with the ~ [200] km s⁻¹ stars belonging to different classes of bar orbits from model for galactic bars as put forth by previous investigations (e.g., Li et al. 2014; Zhou et al. 2017; McGough et al. 2020; Zhou et al. 2021). They are also inconsistent with the idea that the MW harbors a kiloparsec nuclear feature (Debattista et al. 2015, 2018). We therefore currently have no single model that can explain the high-V_{GSR} observations.

To corroborate the accuracy of the Gaia EDR3 proper motions on the plane of the Galaxy, APOGEE stars in a number of globular clusters are identified without the aid of proper motion criteria. Many of these stars had previously not been reported. The first APOGEE stars in the globular clusters 2MASS-GC02 and Terzan 4 are presented, both of which show evidence of multiple populations. The average radial velocity of the stars in 2MASS-GC02 is 87 ± 7 km s⁻¹, which is more than 150 km s⁻¹ offset from the literature value (Borissova et al. 2007).

We thank the referee for suggestions that improved the quality of our analysis. A.M.K and R.E.C. are thankful to colleagues Danielle Miller and Andrew Schulz for aiding with the AAT observations and for stimulating conversations. A.M. acknowledges support from grant AST-2009836 from the National Science Foundation. The grant support provided, in part, by the M.J. Murdock Charitable Trust (NS-2017321) is acknowledged. V.P.D. is supported by STFC Consolidated grant #ST/R000786/1. A.J.K.H gratefully acknowledges funding by the Deutsche Forschungsgemeinschaft (DFG, German Research Foundation)—Project-ID 138713538—SFB 881 (“The Milky Way System”), subprojects A03, A05, A11.

References

AAO software Team 2015, 2ndinf: Data reduction software, Astrophysics Source Code Library, ascl:1505.015
Ahumada, R., Allende Prieto, C., Almeida, A., et al. 2020, ApJS, 249, 3
Alonso-García, J., Saito, R. K., Hempel, M., et al. 2018, A&A, 619, 4
Alonso-García, J., Debattista, V. P., Erwin, P., Kayser, A., Hilker, M., Grebel, E. K., & Willemsen, P. G. 2008, A&A, 474, 121
Carretta, E., Bragaglia, A., Gratton, R., & Lucatello, S. 2009, A&A, 505, 139
Clarke, J. P., Wegg, C., Gerhard, O., et al. 2019, MNRAS, 489, 3519
Debattista, V. P., Erwin, P., Kayser, A., Hilker, M., Grebel, E. K., & Willemsen, P. G. 2008, A&A, 474, 121
Franco, P., & Röser, S. 2015, MNRAS, 454, 3352
McGough, D. P., Evans, N. W., & Sanders, J. L. 2020, MNRAS, 493, 2676
Rosen, J., & Helmi, A. 2007, A&A, 472, 881
Vanzi, L. 2000, A&A, 362, 1
Vanzi, L. 2000, A&A, 362, 1
Vanzi, L. 2000, A&A, 362, 1
Vanzi, L. 2000, A&A, 362, 1
Mészáros, S., Masseron, T., & García-Hernández, D. A. 2020, 
MNRAS, 492, 1641

Minchev, I. 2016, 
AN, 337, 703

Minniti, D., Lucas, P. W., Emerson, J. P., et al. 2010, 
NewA, 15, 433

Molloy, M., Smith, M. C., Evans, N. W., & Shen, J. 2015, 
ApJ, 812, 146

Nandakumar, G., Schultheis, M., Hayden, M., et al. 2017, 
A&A, 606, 97

Nataf, D. M., Gonzalez, O. A., Casagranda, L., et al. 2016, 
MNRAS, 456, 2692

Nataf, D. M., Wyse, R., Schiavon, R. P., et al. 2019, 
AJ, 158, 1

Nidever, D. L., Żasowski, G., Majewski, S. R., et al. 2012, 
ApJ, 755, L25

Nishiya, S., Tamura, M., Hatanaka Kato, D., et al. 2009, 
ApJ, 696, 1407

Nishiya, S., Yasui, K., Nagata, T., Yoshikawa, T., Uchiyama, H., et al. 2013, 
ApJ, 769, 28

Nogueras-Lara, F., Schödel, R., Gallego-Calvente, A. T., et al. 2019, 
A&A, 630, 3

Nogueras-Lara, F., Schödel, R., Gallego-Calvente, A. T., et al. 2020, 
NatAs, 4, 377

Peñaloza, F., Pessev, P., Vañquez, S., et al. 2015, 
PASP, 127, 329

Pérez-Villegas, A., Barbuy, B., Kerber, L. O., et al. 2020, 
MNRAS, 491, 3251

Pietrinferni, A., Cassisi, S., Salaris, M., & Castelli, F. 2004, 
ApJ, 612, 167

Pietrinferni, A., Cassisi, S., Salaris, M., & Castelli, F. 2006, 
ApJ, 642, 797

Portail, M., Wegg, C., Gerhard, O., & Martínez-Valpuesta, I. 2015, 
MNRAS, 448, 713

Queiroz, A. B. A., Anders, F., Chiappini, C., et al. 2020, 
A&A, 638, 76

Queiroz, A. B. A., Chiappini, C., Perez-Villegas, A., et al. 2021, 
arxiv:2007.12915

Recio-Blanco, A., Rojas-Arriagada, A., de Laverny, P., et al. 2017, 
A&A, 602, L14

Renzi, A., D’Antona, F., Cassisi, S., et al. 2015, 
MNRAS, 454, 4197

Rich, R. M., Johnson, C. I., Young, M., et al. 2020, 
MNRAS, 499, 2340

Rojas-Arriagada, A., Zasowski, G., Schultheis, M., et al. 2020, 
MNRAS, 499, 1037

Rossi, L. J., Ortolani, S., Barbuy, B., Bica, E., & Bonfanti, A. 2015, 
MNRAS, 450, 3270

Salinas, R., & Strader, J. 2015, 
ApJ, 809, 169

Sanders, J. L., Smith, L. W., & Evans, N. W. 2019, 
MNRAS, 488, 4552

Schiavon, R. P., et al. 2017, 
MNRAS, 465, 501

Shen, J., Rich, M. R., Kormendy, J., et al. 2010, 
ApJl, 720, L72

Smith, L. C., Lucas, P. W., Kurtz, R., et al. 2018, 
MNRAS, 474, 1826

Tang, B., Cohen, R. E., Geisler, D., et al. 2017, 
MNRAS, 465, 19

Trapp, A. C., Rich, R. M., Morris, M. R., et al. 2018, 
ApJ, 861, 75

VandenBerg, D. A., Brogaard, K., Leaman, R., & Casagrande, L. 2013, 
ApJ, 775, 134

Vasiliev, E. 2019a, 
MNRAS, 484, 2832

Vasiliev, E. 2019b, 
MNRAS, 489, 623

Vasiliev, E., & Baumgardt, H. 2021, 
MNRAS, 505, 5978

Zasowski, G., Ness, M. K., García Pérez, A. E., et al. 2016, 
ApJ, 832, 132

Zasowski, G., Cohen, R. E., Chojnowski, S. D., et al. 2017, 
AJ, 854, 198

Zhou, Y., Shen, J., Liu, C., et al. 2017, 
ApJ, 847, 74

Zhou, Y., Liu, Z. Y., Simion, I., et al. 2021, 
ApJ, 908, 21