We demonstrate chaotic behavior of timelike, null and spacelike geodesics in non-homogeneous vacuum pp-wave solutions. This seems to be the first known example of a chaotic motion in exact radiative spacetime.

In the context of general relativity the first system for which a chaotic behavior of solutions to the Einstein equations has been recognized and thoroughly studied were Bianchi IX cosmological models (see for example [1], [3], and references therein). Complicated nonlinear effects also occur in systems with coupled gravitational and scalar fields [3], [6].

Another type of problems providing nonlinear dynamical systems in general relativity are the studies of geodesic motion in given spacetimes. In particular, chaotic behavior of geodesics in the relativistic analogue of the two fixed-centres problem (modelled by extreme black holes) was examined in [3] - [10]. Chaotic geodesic motion was also found in Schwarzschild spacetime [11] and in topologically non-trivial Robertson-Walker universe [15], [16] and in some static axisymmetric spacetimes [15], [16] and in various other cases (for example in the n-saddle case). This particular case of the Hénon-Heiles Hamiltonian [26] which has been investigated by Rod [22].

Here we investigate motion in exact gravitational waves, namely in the widely known class of vacuum pp-waves [19], the metric of which can be written in the form

$$ds^2 = 2 d\zeta d\bar{\zeta} - 2 du dv - (f + \bar{f}) du^2,$$

(1)

where $f(u, \zeta)$ is an arbitrary function of $u$ and the complex coordinate $\zeta$ spanning the plane wave surfaces $u = \text{const.}$. When $f$ is linear in $\zeta$, the metric (1) represents Minkowski universe. The case $f = g(u)\zeta^2$ describes plane gravitational waves ("homogeneous" pp-waves) which have thoroughly been investigated (see [19] for Refs). This simple example of an exact radiative spacetime has also been used for construction of sandwich and impulsive waves [23], [24].

However, here we wish to study geodesics in more general, non-homogeneous vacuum pp-waves and demonstrate their chaotic behavior. The geodesic equations for (1) are

$$\ddot{\zeta} + \frac{1}{2} \bar{f}\zeta U^2 = 0,$$  

(2)

$$\dot{u} = U = \text{const.},$$  

(3)

$$\ddot{\bar{v}} + (f\zeta + \bar{f}\zeta) U + \frac{1}{2} (f + \bar{f}) U^2 = 0,$$  

(4)

where dot denotes $d/d\tau$ with $\tau$ being an affine parameter. Assuming also a condition normalizing the four-velocity such that $U^\mu U_\mu = \epsilon$ where $\epsilon = -1, 0, +1$ for timelike, null or spacelike geodesics, we get

$$\ddot{v} = \frac{1}{2} U^{-1} \left[ 2\dot{\zeta}^2 - (f + \bar{f}) U^2 - \epsilon \right].$$  

(5)

We consider $U \neq 0$ (for $U = 0$ Eqs. (2)-(4) can be integrated yielding only some trivial geodesics). By differentiating Eq. (3) and using (4) we immediately obtain (2) which can thus be omitted. Hence it suffices to find solutions of Eq. (2) since $v(\tau)$ can then be obtained by integrating Eq. (3), and $u(\tau) = U\tau + u_0$.

The remaining equation (3) has the same form for timelike, null and spacelike geodesics. Introducing real coordinates $x$ and $y$ by $\zeta = x + iy$ we get a system which, for $f$ independent of $u$, follows from the Hamiltonian

$$H = \frac{1}{2} (p_x^2 + p_y^2) + V(x, y),$$  

(6)

where the potential is $V(x, y) = \frac{1}{2} U^2 \Re f$. For non-homogeneous pp-wave spacetimes given by $f = C\zeta^n, C = \text{const.}, n = 3, 4, \ldots$, the corresponding potential

$$V(x, y) = \frac{1}{2} C U^2 \Re \zeta^n,$$  

(7)

is called ‘n-saddle’. It can be visualized in polar coordinates $\rho, \phi$ where $\zeta = \rho \exp(i\phi)$, in which it takes the form $V(\rho, \phi) = \frac{1}{2} C U^2 \rho^n \cos(n\phi)$.

Now, it was shown previously by Rod, Church and Pecelli in a series of papers [23] - [27] that motion in the Hamiltonian system (3) with the n-saddle potential is chaotic.

Let us first briefly summarize their results for the simplest case $n = 3$. The corresponding potential (after removing an unimportant multiplicative factor by a suitable rescaling of $\tau$)

$$V(x, y) = \frac{1}{6} x^3 - xy^2$$  

(8)

is called a ‘monkey saddle’. Interestingly, this is a special case of famous Hénon-Heiles Hamiltonian [26] which is known to be a ‘textbook’ example of a chaotic system (however, their quadratic terms are absent in our case). This particular case of the Hénon-Heiles Hamiltonian has been investigated by Rod [22]. He concentrated on bounded orbits in the energy manifolds...
$H(x, y, p_x, p_y) = E > 0$. The homogeneity of $V$ guarantees that the orbit structure for any two positive values of $E$ is isomorphic modulo a constant scale factor and adjustment of time: $x \to \tilde{x} = \lambda x$, $y \to \tilde{y} = \lambda y$ and $\tau \to \tilde{\tau} = \tau/\sqrt{\lambda}$ results in $E \to \tilde{E} = \lambda^3 E$. Therefore, without loss of generality one can restrict to one particular value of $E$.

In order to describe the topological structure of all bounded orbits Rod first constructed three basic unstable periodic orbits (denoted by $\Pi_j$) which are isolated invariant sets for the flow. The region in which these bounded orbits occur can be decomposed into three disjoint cells $R_j$ (see Ref. 22 or 27 for details); each contain only one orbit $\Pi_j$ and no other bounded orbits. Hence, $\Pi_j$ is the only invariant set in $R_j$, i.e. it is isolated.

Subsequently, Rod investigated orbits asymptotic to basic orbits $\Pi_j$ as $\tau \to \pm \infty$ and showed that these asymptotic sets intersect transversely. This gives the existence of orbits that ‘connect’ the orbits $\Pi_j$: they are homoclinic (asymptotic to the same periodic orbit in both time directions) or heteroclinic (asymptotic to two different periodic orbits, one in each time direction). It is the existence of these orbits that illustrates complicated chaotic structure of the flow.

The topology of possible orbits in phase space can also be represented by symbolic dynamics given here by a set of biinfinite sequences, $s \equiv \cdots, s_k, s_{k+1}, s_{k+2}, \cdots$, where $s_k \in \{1, 2, 3\}$, $s_k \neq s_{k+1}$. Using a topological version of the Smale horseshoe map, it was shown in Ref. 22 that to any biinfinite sequence of symbols $\{1, 2, 3\}$ there exists an uncountable number of bounded orbits running through the blocks $R_j$ in the prescribed order as $\tau$ goes from $-\infty$ to $+\infty$. Also, the flow admits at least a countable number of non-degenerate homoclinic and heteroclinic orbits.

Rod remarked that these results could be refined if the unstable periodic orbits $\Pi_j$ were known to be hyperbolic so that they would admit stable and unstable asymptotic manifolds. Consequently, to each periodic symbol sequence there would correspond a countable collection of periodic orbits. The difficulties in proving the hyperbolicity of $\Pi_j$ were subsequently overcome in Ref. 23. In Ref. 24, summarizing and generalizing some previous results [25], the Hamiltonian \( H(x, y, p_x, p_y) = E > 0 \) was presented as an example of a system for which the Smale horseshoe map can explicitly be embedded as a subsystem into the flow along the homoclinic and heteroclinic orbits. This completed the proof of chaotic behavior of the studied system.

Similar results hold for geodesic motion in arbitrary non-homogeneous $pp$-waves with the structural function of the form $f = C^{\alpha}$, where $\alpha \geq 4$, i.e. for a general $n$-saddle potential \( \Phi \). It was shown in Ref. 22, 25 that the decomposition into isolating cells $R_j$, $j = 1, 2, \cdots, n$, each containing exactly one of the basic unstable periodic solutions $\Pi_j$, is analogous to the case $n = 3$. Subsequently, the orbits $\Pi_j$ were proven to be hyperbolic \( \Pi_j \) and the existence of homoclinic and heteroclinic orbits was established [24]. Again, given any biinfinite sequence, uncountably many orbits can be found which pass from one block containing $\Pi_j$ to the other in the specified order.

In order to independently support these arguments for chaotic behavior of geodesics in non-homogeneous $pp$-waves we investigate the motion also by a fractal method. Complementary to the analysis described above, we concentrate on unbounded geodesics and we do not restrict to the same energy manifold $E = \text{const}$.

Chaos is usually indicated by a sensitive dependence of the evolution on the choice of initial conditions. The coordinate independent fractal method (see for example Ref. 2, 3, 8, 11) starts with a definition of different asymptotic outcomes (given here by ‘types of ends’ of all possible trajectories). A set of initial conditions is evolved numerically until one of the outcome states is reached. Chaos is uncovered if the basin boundaries which separate initial conditions leading to different outcomes are fractal. Such fractal partitions are the result of chaotic dynamics. As we shall now demonstrate, we observe exactly these structures in the studied system.

We integrate numerically the equations of motion given by (1), (2). The initial conditions are chosen such that the geodesics start at $\tau = 0$ from a unit circle in the $(x, y)$-plane (due to the homogeneity of the $n$-saddle potential all other geodesics can simply be obtained by a suitable rescaling). It is natural to parametrize the initial positions by an angle $\phi \in [-\pi, \pi]$ such that $x(0) = \cos \phi$, $y(0) = \sin \phi$. In Fig. 1 we present typical trajectories of geodesics for $n = 3, 4, 5$, when $\dot{x}(0) = 0 = \dot{y}(0)$. We observe that each unbounded geodesic escapes to infinity (where for $n \geq 3$ the curvature singularity is located) only along one of the $n$ channels in the potential. The axes of these outcome channels are given in polar coordinates by the condition $\cos(n\phi_j) = -1$, $j = 1, \cdots, n$, and represent radial lines ‘of steepest descent’ since $V \rightarrow -\infty$ as $\rho \rightarrow \infty$ most rapidly along them. (For non-zero initial velocities more geodesics prefer one of the channels but the character of motion does not change significantly [27].)

In fact, any unbounded geodesic oscillates around the radial axis $\phi_j = (2j - 1)\pi/n$ of the corresponding $j$-th outcome channel. Introducing $\Delta\phi_j(\tau) = \phi(\tau) - \phi_j$, we find asymptotically that $\rho \approx [(n/2 - 1)/\sqrt{C}]^{2/(2-n)}$ as $\rho \rightarrow \infty$ and

$$\Delta\phi_j(\tau) \approx (\tau - \tau_j)^{\alpha} A \cos[b \ln(\tau - \tau_j)] + B \sin[b \ln(\tau - \tau_j)],$$

where $\alpha = \frac{1}{2}(n+2)/(n-2)$, $b = \frac{\sqrt{n}}{\sqrt{7}}$, $\tau \rightarrow \tau_j$, frequency of their oscillations around $\phi_j$ grows to infinity while the amplitude of oscillations tends to zero. We may call this effect a ‘focusation’. 
Let us return back to our observation that all unbounded geodesics approach infinity through only \( n \) distinct channels. These represent possible outcomes of our system and we assign them symbol \( j \) which takes one of the corresponding values, \( j \in \{1, 2, \ldots, n\} \) (thus, for example, \( j = 1 \) means that the geodesic approach infinity at \( \rho = \infty \) through the first channel with the axis \( \phi_1 = \pi/n \) as \( \tau \to \tau_s > 0 \)). From Fig. 1 we observe that in certain regions the function \( j(\phi) \) depends very sensitively on initial position given by \( \phi \). We calculated \( j(\phi) \) numerically for \( n = 3, 4, 5 \) — the results are shown in Fig. 2. Also, in the same diagrams we plot the function \( \tau_s(\phi) \) which takes the value of the parameter \( \tau \) when the singularity at \( \rho = \infty \) is reached by a given geodesic.

Clearly, the boundaries between the outcomes appear to be fractal which can be confirmed on the enlarged detail of the image and the enlarged detail of the detail etc. In Fig. 3 we show such zooming in of the fractal interval localized around the value \( \phi \approx 0 \) for \( n = 3 \) (there are two symmetric fractal intervals in this case around \( \phi \approx \pm \frac{2}{3} \pi \)) up to the sixth level. At each level the structure has the same property, namely that between two larger connected sets of geodesics with outcome channels \( j_1 \) and \( j_2 \neq j_1 \) there is always a smaller connected set of geodesics with outcome channel \( j_3 \) such that \( j_3 \neq j_1 \) and \( j_3 \neq j_2 \). Similarly as in [7], [10], the structure of initial conditions resembles three mixed Cantor sets, and this fact is a manifestation of chaos.
The fractal structure described by $j(\phi)$ and $\tau_s(\phi)$ is clearly confirmed here by zooming in the interval around $\phi \approx 0$ for $n = 3$.

The above structure of $j(\phi)$ has its counterpart in the fractal structure of $\tau_s(\phi)$, see Fig. 2 and Fig. 3. We observe that the value of $\tau_s$ goes to infinity on each discontinuity of $j(\phi)$, i.e., on any fractal basin boundary between the different outcomes. There is an infinite set of peaks corresponding to chaotic bounded orbits which never ‘decide’ on a particular outcome, and so never escape to infinity. The value of $\tau_s$ also increases in non-chaotic regions of $\phi$ as one zooms in the higher levels of the fractal. This is natural since these levels are given by geodesics which undergo ‘more bounces’ on the potential walls before falling into one of the outcome channels so that their values of $\tau_s$ are greater.

We demonstrated chaotic behavior of geodesics in non-homogeneous $pp$-waves by invariant analytic and numerical methods. As far as we know, this is the first explicit demonstration of chaos in exact radiative space-times (chaotic interaction of particles with linearized gravitational waves on given backgrounds has already been studied in [11], [13], [28] - [30]). Since $pp$-wave solutions are the simplest gravitational waves it would be an interesting task to search for a chaotic motion in other radiative space-times.

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[1] D. Hobill, A. Burd, and A. Coley (eds), Deterministic Chaos in General Relativity (New York, Plenum Press, 1994).
[2] N. J. Cornish and J. J. Levin, Phys. Rev. Lett. 78, 998 (1997); Phys. Rev. D 55, 7489 (1997).
[3] M. W. Choptuik, Phys. Rev. Lett. 70, 9 (1993).
[4] E. Calzetta and C. El Hasi, Classical Quantum Gravity 10, 1825 (1993).
[5] S. Blanco, A. Costa, and O. A. Rosso, Gen. Relativ. Grav. 27, 1295 (1995).
[6] N. J. Cornish and J. J. Levin, Phys. Rev. D 53, 3022 (1996).
[7] G. Contopoulos, Proc. R. Soc. London A 431, 183 (1990); 435, 551 (1991).
[8] C. P. Dettmann, N. E. Frankel, and N. J. Cornish, Phys. Rev. D 50, R618 (1994).
[9] U. Yurtsever, Phys. Rev. D 52, 3176 (1995).
[10] N. J. Cornish and G. W. Gibbons, Classical Quantum Gravity 14, 1865 (1997).
[11] L. Bombelli and E. Calzetta, Classical Quantum Gravity 9, 2573 (1992).
[12] W. M. Vieira and P. S. Letelier, Phys. Rev. Lett. 76, 1409 (1996).
[13] P. S. Letelier and W. M. Vieira, Classical Quantum Gravity 14, 1249 (1997).
[14] S. Suzuki and K. Maeda, Phys. Rev. D 55, 4848 (1997).
[15] V. Karas and D. Vokrouhlický, Gen. Relativ. Grav. 24, 729 (1992).
[16] Y. Sota, S. Suzuki, and K. Maeda, Classical Quantum Gravity 13, 1241 (1996).
[17] C. M. Lockhart, B. Misra, and I. Prigogine, Phys. Rev. D 25, 921 (1982).
[18] R. Tomaszcz, J. Math. Phys. 32, 2571 (1991).
[19] D. Kramer, H. Stephani, M. A. H. MacCallum, and E. Herlt, Exact Solutions of Einstein’s Field Equations (Cambridge, Cambridge University Press, 1980).
[20] H. Bondi, F. Pirani, and I. Robinson, Proc. R. Soc. London A 251, 519 (1959).
[21] R. Penrose, in General Relativity, edited by L. O’Raifeartaigh (Clarendon Press, Oxford, 1972).
[22] D. L. Rod, J. Differential Equations 14, 129 (1973).
[23] D. L. Rod, G. Peceli, and R. C. Churchill, J. Differential Equations 24, 329 (1977).
[24] R. C. Churchill and D. L. Rod, J. Differential Equations 37, 23 (1980).
[25] R. C. Churchill and D. L. Rod, J. Differential Equations 21, 39 (1976); 21, 66 (1976).
[26] M. Hénon and C. Heiles, Astronom. J. 69, 73 (1964).
[27] J. Podolský and K. Veselý, in preparation (1998).
[28] H. Varvoglis and D. Papadopoulos, Astron. Astrophys. 261, 664 (1992).
[29] C. Chicone, B. Mashhoon, and D. G. Retzloff, Classical Quantum Gravity 14, 699 (1997).
[30] F. Kokubun, Phys. Rev. D 57, 2610 (1998).
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