MAGIC and Fermi-LAT gamma-ray results on unassociated HAWC sources

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J. P. Harding,26 S. Hernandez,27 F. Hueyotl-Zahuantitla,28 C. M. Hui,42
P. Hüntemeyer,41 A. Irriarte,32 A. Jardin-Blicq,43 V. Joshi,43 S. Kaufmann,28 A. Lara,40
R. J. Lauer,44 W. H. Lee,32 D. Lennarz,45 H. León Vargas,27 J. T. Linnemann,46
A. L. Longinotti,34 G. Luis-Raya,47 R. Luna-García,48 R. López-Coto,43 K. Malone,30
S. S. Marinelli,46 O. Martinez,36 I. Martínez-Castellanos,39 J. Martínez-Castro,48
H. Martínez-Huerta,49 J. A. Matthews,44 P. Miranda-Romagnoli,50 E. Moreno,36
M. Mostafá, A. Nayerhoda,35 L. Nellen,51 M. Newbold,52 M. U. Nisa,31
R. Noriega-Papaqui,50 R. Pelayo,48 J. Pretz,30 E. G. Pérez-Pérez,47 Z. Ren,44
C. D. Rho,31 C. Rivière,39 D. Rosa-González,34 M. Rosenberg,30 E. Ruiz-Velasco,43
F. Salesa Greus,35 A. Sandoval,27 M. Schneider,53 M. Seglar Arroyo,30 G. Sinnis,26
A. J. Smith,39 R. W. Springer,52 P. Surajbali,43 I. Taboada,45 O. Tibolla,28
K. Tollefson,46 I. Torres,34 T. N. Ukwatta,26 G. Vianello,54 L. Villaseñor,36 F. Werner,43
S. Westerhoff,33 J. Wood,33 T. Yapici,31 G. Yodh,55 A. Zepeda,49 H. Zhou,26
J. D. Álvarez,29 HAWC Collaboration, M. Ajello,56 L. Baldini,57 G. Barbiellini,58,59
B. Berenji,60 E. Bissaldi,61,62 R. D. Blandford,63 R. Bonino,64,65 E. Bottacini,63,66
T. J. Brandt,67 J. Bregeon,68 P. Bruel,69 R. A. Cameron,63 R. Caputo,70 P. A. Caraveo,71
D. Castro,67,72 E. Cavazzuti,73 G. Chiaro,71 S. Ciprini,73,74 D. Costantin,76
F. D'Ammando,77,78 F. de Palma,62,79 A. Desai,56 N. Di Lalla,57 M. Di Mauro,63
L. Di Venere,61,62 A. Domínguez,80 C. Favuzzi,61,62 Y. Fukazawa,81 S. Funk,82
P. Fusco,61,62 F. Gargano,62 D. Gasparrini,74,75 N. Giglietto,61,62 F. Giordano,61,62
M. Giroletti,77 T. Glanzman,63 D. Green,67,68 I. A. Grenier,84 S. Guiriec,85,67
A. K. Harding,67 E. Hays,67 J. W. Hewitt,86 D. Horan,69 G. Jóhannesson,87,88
M. Kuss,89 S. Larsson,90,91 I. Liodakis,63 F. Longo,58,59 F. Loparco,61,62 P. Lubrano,75 J.
D. Magill,83 S. Maldera,64 A. Manfreda,57 M. N. Mazziotta,62 I. Mereu,92 P.
F. Michelson,63 T. Mizuno,93 M. E. Monzani,63 A. Morselli,94 I. V. Moskalenko,63
M. Negro,64,65 E. Nuss,68 N. Omodei,63 M. Orienti,77 E. Orlando,63 J. F. Ormes,95
M. Palatiello,57,58 V. S. Paliya,56 M. Persic,58,96 M. Pesce-Rollins,89 V. Petrosian,63
F. Piron,68 T. A. Porter,63 G. Principe,82 S. Rainò,61,62 B. Rani,67 M. Razzano,89
S. RAZZANO,97 A. REIMER,63,98 O. REIMER,63,98 C. SGRÒ,89 E. J. SISKIND,99 G. SPANDRE,89
P. Spinelli,60,61 H. Tajima,63,100 M. Takahashi,101 J. B. Thayer,63 D. J. Thompson,67 D.
R. Torres,102,103 E. TORRESI,104 E. Troja,66,82 J. Valverde,69 K. Wood,105 and
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ABSTRACT

The HAWC Collaboration released the 2HWC catalogue of TeV sources, in which 19
show no association with any known high-energy (HE; $E \gtrsim 10$ GeV) or very-high-energy
(VHE; $E \gtrsim 300$ GeV) sources. This catalogue motivated follow-up studies by both the
Major Atmospheric Gamma-ray Imaging Cherenkov (MAGIC) and Fermi-LAT (Large Area Telescope) observatories with the aim of investigating gamma-ray emission over a broad energy band. In this paper, we report the results from the first joint work between High Altitude Water Cherenkov (HAWC), MAGIC, and Fermi-LAT on three unassociated HAWC sources: 2HWC J2006$+341$, 2HWC J1907$+084^*$, and 2HWC J1852$+013^*$. Although no significant detection was found in the HE and VHE regimes, this investigation shows that a minimum 1° extension (at 95 per cent confidence level) and harder spectrum in the GeV than the one extrapolated from HAWC results are required in the case of 2HWC J1852$+013^*$, whilst a simply minimum extension of 0.16° (at 95 per cent confidence level) can already explain the scenario proposed by HAWC for the remaining sources. Moreover, the hypothesis that these sources are pulsar wind nebulae is also investigated in detail.

**Key words:** pulsars: general – gamma rays: general – astroparticle physics.

# 1 INTRODUCTION

The synergy of current gamma-ray observatories is a unique opportunity to explore the Universe between a few tens of MeV to hundreds of TeV. The study of such a broad energy band can be accomplished by means of different detection techniques. At the highest energies are the Water Cherenkov arrays, such as the High Altitude Water Cherenkov (HAWC) Observatory, which is sensitive to cosmic and gamma rays between 100 GeV and 100 TeV. A large effective area and high duty cycle make HAWC an optimal instrument to perform survey studies on multi-TeV sources. Other detection techniques achieve a better sensitivity than HAWC at energies below approximately 10 TeV: the Imaging Atmospheric Cherenkov Telescope (IACT) technique, as implemented by the Major Atmospheric Gamma-ray Imaging Cherenkov (MAGIC) telescopes, provides better angular and energy resolution. However, the duty cycle of such atmospheric Cherenkov telescopes is greatly reduced by high-intensity background light or non-optimal weather conditions at night. At even lower energies (∼GeV), detectors on board satellites, like the Fermi-Large Area Telescope (LAT), are used to detect gamma rays. This instrument has a high duty cycle, a wide field of view (FoV) of 20 per cent of the sky and very good gamma/hadron separation.

HAWC has published two catalogues of TeV sources: 1HWC for sources in the inner Galactic plane using 275 d of data with a configuration of approximately one-third of the full array (HAWC-111; Abeysekara et al. 2016), and 2HWC for almost the entire sky using 507 d of the completed HAWC detector (Abeysekara et al. 2017c). The second catalogue improves over the first with respect to exposure time, detector size, and angular resolution, resulting in a significant improvement in sensitivity. As done for the previous catalogue, 2HWC data were analysed using a binned likelihood method described in Younk et al. (2015). In this method, a source model needs to be assumed for all sources in the sky. The model for each source is characterized by the source morphology and its spectrum. For the 2HWC analysis, HAWC used two different approaches: (1) a point-like search adopting a spectrum defined by a power-law function, \( dN/dE = N_0 \left( \frac{E}{E_0} \right)^{-\Gamma} \) (with \( N_0 \) the normalization, \( E_0 \) the pivot energy and \( \Gamma \) the spectral index), with spectral index \( \Gamma = 2.7 \), and (2) extended source searches with a source morphology modelled as a uniform disc of 0.5°, 1°, and 2° in radius and spectral index \( \Gamma = 2.0 \). The total number of sources identified in this catalogue was 39, of which 19 were not associated with any previously reported TeV source within an angular distance of 0.5°. All 2HWC sources presented a test statistic (TS) above 25 (equivalent to a pre-trial significance of ∼5σ).

The 2HWC catalogue motivated follow-up studies with H.E.S.S. (López-Coto et al. 2017), VERITAS (Abeysekara et al. 2018), and also MAGIC and Fermi-LAT. In this work, we focused on the 19 sources with no high-energy (HE; \( E \gtrsim 10 \) GeV) or very-high-energy (VHE; \( E \gtrsim 300 \) GeV) association in order to provide new multiwavelength information of candidates without a lower energy counterpart. After evaluating those sources, a short list of three targets was selected: 2HWC J2006$+341$ (RA = 301.55°, Dec. = 34.18°), 2HWC J1907$+084^*$ (RA = 286.79°, Dec. = 8.50°), and 2HWC J1852$+013^*$ (RA = 283.01°, Dec. = 1.38°). These sources were chosen because they lie in the FoV of previous MAGIC observations, allowing MAGIC to analyse these sources without performing new dedicated observations.

Even though the HAWC spectra of each source were determined using a likelihood fit, the 2HWC catalogue did not use a likelihood method to describe multiple sources simultaneously. In the 2HWC catalogue, the asterisk of 2HWC J1907$+084^*$ and 2HWC J1852$+013^*$ indicates that the sources were near another source with larger significance and thus their characterization may be influenced by neighbouring sources. 2HWC J2006$+340$, 2HWC J1907$+084^*$, and 2HWC J1852$+013^*$ were detected in the point source search with significances of 6.10σ, 5.80σ, and 8.50σ, respectively. The corresponding photon index and flux normalization values obtained in the 2HWC catalogue are listed in Table 1. Their corresponding energy range is computed with a dedicated HAWC analysis (see Section 2.1) and also given in the table.

The paper is structured as follows. In Section 2, the data analysis for MAGIC and Fermi-LAT observations are presented. Description of the specific HAWC analysis on the selected sources is also included. The observations and results, for each source separately, are shown in Section 3. Discussion and conclusion can be found in Sections 4 and 5, respectively.

# 2 DATA ANALYSIS

## 2.1 HAWC

The HAWC Observatory is the second generation of ground-based gamma-ray extensive air shower arrays, located in Sierra Negra, Mexico (19.0° N, 97.3° W, 4100 m a.s.l.), and successor to the Milagro Gamma-ray Observatory. The current system, inaugurated on 2015 March 20, is comprised of 300 water Cherenkov detectors over an area of 22 000 m². Science operations began before
Table 1. Coordinates, photon index, flux at the pivot energy (7 TeV), and energy range for the analysed sources. Values are provided in the 2HWC catalogue, except for the energy range, which was obtained in a dedicated analysis. Only statistical uncertainties are shown. Based on a study of the Crab Nebula by HAWC (Abeysekara et al. 2017b), the systematic uncertainty can be divided into several components: 0.10° in angular resolution, 0.2° in photon index, and 50 per cent in flux normalization.

| RA (°) | Dec. (°) | α (°) | b (°) | 1σ stat. error | Photon index | Flux normalization (×10⁻¹⁵ TeV⁻¹ cm⁻² s⁻¹) | Energy range (TeV) |
|-------|---------|-------|-------|----------------|--------------|------------------------------------------|------------------|
| 2HWC J2006+341 | 301.55 | 34.18 | 71.33 | 1.16 | 0.13 | 2.64 ± 0.15 | 9.6 ± 1.9 | 1–86 |
| 2HWC J1907+084* | 286.79 | 8.50 | 42.28 | 0.14 | 0.27 | 3.25 ± 0.18 | 7.3 ± 2.5 | 0.18–10 |
| 2HWC J1852+013* | 283.01 | 1.38 | 34.23 | 0.50 | 0.13 | 2.90 ± 0.10 | 18.2 ± 2.3 | 0.4–50 |

detector completion, under the HAWC-111 configuration. The angular resolution of HAWC varies with event size (fraction of photomultiplier tubes reporting a signal or $f_{\text{hit}}$) from 0.17° to 1.0° (Abeysekara et al. 2017b). HAWC operates with >95 per cent duty cycle with a large FoV of 15 per cent of the sky, which allows it to scan two-thirds of the sky every 24 h.

Information presented here on 2HWC J2006+341, 2HWC J1907+084*, and 2HWC J1852+013* is taken mostly from the 2HWC catalogue Abeysekara et al. (2017c). The only exception is the energy range shown in Table 1. The likelihood analysis in HAWC is computed over $f_{\text{hit}}$ bins, which can be considered an energy estimator. However, the $f_{\text{hit}}$ bins depend strongly on the declination and spectral hardness of each source, and so does this $f_{\text{hit}}$/energy correlation. The energy range is then given as a constraint on the photon distribution as a function of $f_{\text{hit}}$ for each separate source. Following Abeysekara et al. (2017c), we take the energy range as the boundaries within which the events contribute to the 75 per cent of the TS value.

2.2 MAGIC

MAGIC is a stereoscopic system of two 17 m diameter IACTs situated on the Canary island of La Palma, Spain (28.8° N, 17.8° W, 2225 m a.s.l.). The current system achieves an integral sensitivity of 0.66 ± 0.03 per cent of the Crab Nebula flux (CU) in 50 h of observation above 220 GeV (Aleksić et al. 2016). The energy threshold in stereoscopic mode is as low as 50 GeV at low zenith angles under dark observational conditions (Aleksić et al. 2012).

The analysis presented in this work is performed using the standard MAGIC analysis software (MARS; Zanin et al. 2013). The significance is computed following equation 17 of Li & Ma (1983). Differential and integral flux upper limits (ULs) are calculated using the Rolke algorithm (Rolke, López & Conrad 2005) with a confidence level (CL) of 95 per cent, assuming a Poissonian background and a total systematic uncertainty of 30 per cent.

As mentioned above, the three analysed sources were included in the FoV of former MAGIC observations. These archival data were taken using the false-source tracking mode, or wobble-mode: the telescopes point at four different positions located 0.4° from the nominal source, which allows us to evaluate the background simultaneously (Fomin et al. 1994). Thus, our observations were not dedicated to the 2HWC sources and so, their coordinates are shifted from the camera centre by different distances than the standard offset of 0.4° (see Fig. 1). To account for their location in the camera, the background used in the calculation of ULs was evaluated through the off-from-wobble-partner (OfWP) method (Zanin et al. 2013).

Table 2 summarizes the distances between the camera centre and the 2HWC sources at the four different wobble positions. The total observation time, after data quality cuts, for each case is also quoted in Table 2. It is worth highlighting that the MAGIC sensitivity depends on the angular offset from the pointing direction. However, after the MAGIC upgrade of 2011–2012, the sensitivity at offset angles larger than 0.4° improved considerably as shown by Aleksić et al. (2016). For the analysis performed in this work, and given the range of angular offsets for all the candidates, the sensitivity remains between ~0.6 and 1.0 per cent CU.

Observations of 2HWC J1852+013* were carried out entirely under dark conditions, i.e. in absence of moonlight. On the other hand, 2HWC J2006+341 and 2HWC J1907+084* were observed with nominal high-voltage at background levels ranging between 1 and 8 times the brightness of the dark sky due to different Moon phases. The higher the moonlight level, the brighter is the night sky background and therefore, stronger cuts to the signal are applied during this analysis, following the prescription of Ahnen et al. (2017). This is taken into account by selecting appropriate Monte Carlo-simulated gamma-ray and background data to match the observational conditions. The background data are used for the computation of the gamma/hadron separation through the Random Forest, a multidimensional classification algorithm based on decision trees (Albert et al. 2008).

Since the three HAWC sources each have a maximum significance in the point-source HAWC maps, they may be point-like sources for MAGIC as well. Therefore, we analyse the candidates under two hypotheses: we assume that the sources are point-like for MAGIC (point spread function (PSF) ≲ 0.10°, beyond a few hundred GeV) or are extended with a radius of 0.16°. Larger extensions cannot be adopted due to the OfWP method and the standard 0.4° offset applied in the wobble pointing mode, because the expected region of gamma-ray emission from the 2HWC source and the background regions selected to compute flux ULs would overlap.

2.3 Fermi-LAT

The LAT on board the Fermi Gamma-Ray Space Telescope has continuously monitored the sky since 2008. It is sensitive to HE gamma rays between 20 MeV and ~1 TeV (Atwood et al. 2009) and scans the entire sky every 3 h. For this work, we used data taken between 2008 August and 2017 February based on the Pass 8 SOURCE photon reconstruction. The Pass 8 data offer two primary benefits for the study of HE gamma-ray sources: a greater...
acceptance compared with previous LAT reconstructions and an improved PSF with a 68 per cent containment angle less than 0.2° above 10 GeV that is nearly constant with increasing energy (Atwood et al. 2013).

For each source of interest, we analyse energies between 10 GeV and 1 TeV using the standard binned likelihood framework provided by the Fermi Science Tools (v10r01p01). Data within a 10° radius of interest were binned into 8 energy bins per decade and a spatial bin size of 0.05°. We used the recommended Galactic and isotropic backgrounds.1

We did not extend our analysis to lower energies for two primary reasons: for any HAWC source to be detected at lower energies, it must be detectable at >10 GeV with the LAT, unless the HAWC and LAT emission is produced by a different component; and this high-energy cut suppresses photons from gamma-ray pulsars and Galactic diffuse emission in the plane. As a source model for this analysis, we use the Third Catalog of Hard Fermi-LAT Sources for point sources (3FHL; Ajello et al. 2017) and the Fermi Galactic Extended Source (FGES) catalogue for extended sources (Ackermann et al. 2017).

Using ~8.5 yr of Pass 8 data, we search for new sources separately testing both a point-like source and an extended source at the location of the HAWC candidate. The spectrum of the source is modelled as a simple power law. After initially fitting a putative point source at the HAWC position, the position is re-localized. The normalizations of other sources within 5° are left as free parameters in the fit. To search for a possible extended source, we use a uniformly illuminated disc with a radius of 0.2° as the initial spatial model. The fermipy package (Wood et al. 2017) fits both the radius and centroid of the possible extended source. If a statistically significant source is not found, Uls at the position of the HAWC source are computed at 95 per cent CL using a Bayesian method. The assumed spectral indices are 2.0, 3.0, and the index

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1Galactic interstellar emission model: gll_jem8_v06.fits, Isotropic: iso_P8R2_SOURCE_V6_v06.txt. Please see http://fermi.gsfc.nasa.gov/ssc/data/access/lat/BackgroundModels.html.
reported in the 2HWC catalogue (see Table 1). Extended source ULs are computed assuming a radius of 0.16° for comparison with the limits placed by MAGIC.

3 OBSERVATIONS AND RESULTS

In the following section, we describe the different regions of the sky that contain the three selected sources, along with the corresponding observations and results. MAGIC differential ULs are listed in Table 3, whilst Fermi-LAT integral ULs (above 10 GeV) are quoted in Table 4 for both point-like and extended hypotheses. Fig. 1 presents the MAGIC significance skymaps for 2HWC J2006+341 and 2HWC J1907+084° assuming an extended source with a 0.16° radius. A smaller 1° × 1° MAGIC significance skymap centred in 2HWC J1852+013° is shown in Fig. 2. The skymap of the entire FoV for this source will be included in a dedicated MAGIC paper on the surrounding region that is in preparation. The flat significance field displayed in all skymaps is compatible with background in the entire FoV. The multiwavelength spectral energy distribution (SED) for each 2HWC source is presented in Fig. 3.

3.1 2HWC J2006+341

2HWC J2006+341 is in the FoV (at ~0.63°) of the compact radio/optical nebula G70.7+1.2, which is thought to be powered by a pulsar–binary system interacting with a surrounding molecular cloud. An unidentified source, 3FHL J2004.2+3339, was detected at the position of this putative binary system. Therefore, VHE gamma-ray emission from the G70.7+1.2 region could be expected due to the interaction between the pulsar wind with both the stellar wind of the companion star and the molecular cloud.

MAGIC observed 3FHL J2004.2+3339 with an extended range of zenith angles from 5° to 50°. The total data sample amounts to ~61 h of good quality data from 2015 April to 2016 August. No significant signal is found in the direction of 2HWC J2006+341 as either a point-like or extended source. In order to calculate the integral flux ULs, MAGIC adopts a power-law distribution with photon index $\Gamma = 2.64$, following the HAWC results. Under point-like assumption, the integral UL, computed at 95 per cent CL for $E > 300$ GeV and index $\Gamma = 3.25$ is $2.8 \times 10^{-13}$ photons cm$^{-2}$ s$^{-1}$ for the point-like and extended hypotheses, respectively. MAGIC does not find any significant excess at the position of 1HWC J1904+080c either, which leads to an integral flux UL for energies greater than 300 GeV of $4.1 \times 10^{-12}$ photons cm$^{-2}$ s$^{-1}$, assuming a power-law spectrum of $\Gamma = 2.6$. From the Fermi-LAT, the $\text{P}_{\text{P}}$ 8 analysis yields no significant emission in the direction of 2HWC J1907+084°, either during the point-like or the extended analysis.

3.3 2HWC J1852+013°

2HWC J1852+013° is located in the FoV of the W44 SNR, as well as in the FoV of the established VHE sources HESS J1858+020 and HESS J1857+026 (subdivided in two emission sites MAGIC J1857.2+0263 and MAGIC J1857.6+0297; MAGIC Collaboration et al. 2014). The region was thus extensively observed by the MAGIC collaboration. 2HWC J1852+013° is also located at 0.56° away from 3FGL J1852.8+0158, which is classified as a probable young pulsar using machine learning techniques (Saz Parkinson et al. 2016).

The data set used by MAGIC here comprises approximately 120 h of dark quality data, taken from 2013 April to 2014 June, with a span in zenith range from 25° to 50°. MAGIC does not find any excess in the direction of 2HWC J1852+013°. Adopting $\Gamma = 2.90$, the constraining 95 per cent CL integral ULs are $3.8 \times 10^{-13}$ photons cm$^{-2}$ s$^{-1}$ for the point-like search and $1.7 \times 10^{-13}$ photons cm$^{-2}$ s$^{-1}$ for the extended search. Specific background selection using OWF was applied in this case, ensuring that no background control region overlaps with any of the several VHE emitting sources in the FoV. As per the previous sources, ULs are given for $E > 300$ GeV. Neither catalogued nor new sources from the $\text{P}_{\text{P}}$ 8 analysis arises in the Fermi-LAT analysis of 2HWC J1852+013°.

4 DISCUSSION

Given that the largest population of TeV emitters in our Galaxy are pulsar wind nebulae (PWNe; see e.g. H.E.S.S. Collaboration et al. 2018a), the selected candidates may be expected to be this source type. However, the lack of a detection by either MAGIC or Fermi-LAT complicates the identification of these sources. In order to investigate their possible PWN nature, we look for detected pulsars near these 2HWC sources using the ATNF catalogue2 (Manchester et al. 2005). According to the characteristic ages of the pulsars around the three selected 2HWC sources (all above a few tens of kyr), if these pulsars had high initial kick velocities, they could now be significantly offset from their initial positions and have left behind an old PWN with no compact object powering it. In this case, the pulsar position is shifted from the PWN, and without injection of magnetic flux, the nebula’s emission is expected to be dominated by inverse Compton (IC).

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2http://www.atnf.csiro.au/people/pulsar/psrcat/

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could only be explained with an improbably large kick velocity for 2HWC J2006+341, 2HWC J1907+084\*, and 2HWC J1852+013\*, assuming a power-law spectrum with spectral index of $\Gamma = 2.64$, 3.25, and 2.90, respectively. ULs for both point-like ($<0.10$) and extended ($<0.16$ radius) assumptions are shown in each case. Due to low statistics, ULs at the highest energy ranges are not always computed for 2HWC J2006+341 and 2HWC J1907+084\*.

| Energy range [GeV] | 2HWC J2006+341 | 2HWC J1907+084\* (photons cm$^{-2}$ s$^{-1}$) | 2HWC J1852+013\* | 2HWC J1907+084\* |
|-------------------|---------------|---------------------------------|-----------------|-----------------|
| Point-like | Extended | Point-like | Extended | Point-like | Extended |
| 139.2–300.0 | $2.6 \times 10^{-11}$ | $6.2 \times 10^{-11}$ | $7.1 \times 10^{-11}$ | $3.1 \times 10^{-10}$ | $1.7 \times 10^{-11}$ | $4.6 \times 10^{-11}$ |
| 300.0–646.3 | $1.4 \times 10^{-12}$ | $1.0 \times 10^{-11}$ | $8.0 \times 10^{-12}$ | $1.7 \times 10^{-11}$ | $8.6 \times 10^{-13}$ | $4.9 \times 10^{-12}$ |
| 646.3–1392.5 | $2.5 \times 10^{-13}$ | $1.3 \times 10^{-12}$ | $2.5 \times 10^{-12}$ | $2.3 \times 10^{-12}$ | $9.0 \times 10^{-14}$ | $3.9 \times 10^{-13}$ |
| 1392.5–3000.0 | $6.0 \times 10^{-14}$ | $9.9 \times 10^{-14}$ | $1.7 \times 10^{-13}$ | $1.0 \times 10^{-12}$ | $6.7 \times 10^{-14}$ | $1.2 \times 10^{-13}$ |
| 3000.0–6463.3 | $1.8 \times 10^{-14}$ | $2.7 \times 10^{-14}$ | $-\$ | $1.4 \times 10^{-13}$ | $7.6 \times 10^{-15}$ | $1.3 \times 10^{-14}$ |
| 6463.3–13924.8 | $-\$ | $9.5 \times 10^{-15}$ | $-\$ | $1.3 \times 10^{-14}$ | $1.1 \times 10^{-14}$ | $5.5 \times 10^{-14}$ |

Table 4. *Fermi*-LAT 95 per cent CL flux ULs, above 10 GeV, assuming point-like source and extended source a radius of 0.16\'.

| $\Gamma_{\text{HAWC}}$ | $\Gamma = 2$ | $\Gamma = 3$ |
|----------------|-------------|-------------|
| Point-like (\times 10^{-11} photons cm^{-2} s^{-1}) | Extended | Point-like (\times 10^{-11} photons cm^{-2} s^{-1}) | Extended | Point-like (\times 10^{-11} photons cm^{-2} s^{-1}) | Extended |
| J2006+341 | 2.4 | 4.4 | 2.3 | 4.7 | 2.4 | 4.2 |
| J1907+084\* | 3.1 | 3.1 | 2.7 | 2.7 | 3.2 | 3.2 |
| J1852+013\* | 2.1 | 3.7 | 2.0 | 3.3 | 2.0 | 3.7 |

Table 3. MAGIC differential ULs (at 95 per cent CL) for 2HWC J2006+341, 2HWC J1907+084\*, and 2HWC J1852+013\*, assuming a power-law spectrum with spectral index of $\Gamma = 2.64$, 3.25, and 2.90, respectively. ULs for both point-like ($<0.10$) and extended ($<0.16$ radius) assumptions are shown in each case. Due to low statistics, ULs at the highest energy ranges are not always computed for 2HWC J2006+341 and 2HWC J1907+084\*.

PSR J2004+3429 is the closest known pulsar to 2HWC J2006+341 at a separation of 0.40\*, and is the only one within a 1° radius. This pulsar lies at a distance of 10 kpc, displays a spin-down power of $\dot{E} = 5.8 \times 10^{35}$ erg s$^{-1}$, and has a characteristic age of $\tau = 18$ kyr. Although energetic enough to power a TeV PWN (see H.E.S.S. Collaboration et al. 2018b), the distance between 2HWC J2006+341 and PSR J2004+3429 makes this connection improbable: given the characteristic age of 18 kyr, an offset of 0.40\* (~70 pc) could only be explained with an improbably large kick velocity for the pulsar of ~4000 km s$^{-1}$. The mean 2D speed for both young and old (~3 Myr) pulsars was determined to be only 307 ± 47 km s$^{-1}$ by Hobbs et al. (2005) with a study involving a subsample of ~50 pulsars’ proper motion. The offset may be considerably less when considering HAWC systematic and statistical errors on the 2HWC source location of $0.40^\circ \pm 0.10^\circ_{\text{sys}} \pm 0.13^\circ_{\text{stat}}$. Assuming the most constraining possible value, 0.24\*, the necessary kick velocity would decrease to ~2300 km s$^{-1}$. This value is not far away from the fastest known pulsar at ~1500 km s$^{-1}$ (Hobbs et al. 2005), though that value is also uncertain given the distance model applied. The highest speed for a pulsar with a well-measured distance is only 640 km s$^{-1}$. Therefore, we conclude that it is unlikely that PSR J2004+3429 is directly responsible for the emission detected by HAWC. On the other hand, Linden et al. (2017) evaluated the probability of random association between 15 2HWC sources and their nearby pulsars, including 2HWC J2006+341 and PSR J2004+3429. For this case, they obtained a chance overlap of only 8 per cent (assuming a source extension of 0.9\* as provided in the 2HWC catalogue by assigning the halo-like structures visible in the residual skymaps to 2HWC J2006+341, which presents its own uncertainties).

There are two pulsars within 0.50\* of 2HWC J1907+084\*: PSR J1908+0833 at 0.30\* and PSR J1908+0839 at 0.33\*. The former is located at a distance of ~11 kpc, with a characteristic age of $\tau = 4.1$ Myr and a spin-down power of $\dot{E} = 5.8 \times 10^{32}$ erg s$^{-1}$. The very low spin-down power and old age make it unlikely to be currently powering a TeV PWN. Alternatively, PSR J1908+0839, at 8.3 kpc and with a characteristic age of $\tau = 1.2$ Myr, is more energetic with $\dot{E} = 1.5 \times 10^{34}$ erg s$^{-1}$, so a relation between this pulsar and the 2HWC source cannot be initially ruled out. As done for 2HWC J2006+341, we calculate the kick velocity for the pulsar, now with an offset of 0.33\* and a characteristic age of 1.2 Myr. The obtained velocity is ~40 km s$^{-1}$, which is low relative to the average kick velocity observed through proper motion studies but remains to be a valid possibility (see fig. 4b from Hobbs et al. 2005).
Figure 3. SED from 10 GeV up to ~90 TeV. In all cases, the assumed spectrum for the sources is a power-law function with photon index $\Gamma = 2.64$ for 2HWC J2006+341 (top), $\Gamma = 3.25$ for 2HWC J1907+084* (middle) and $\Gamma = 2.90$ for 2HWC J1852+013* (bottom), as obtained by HAWC (see Table 1). Fermi-LAT 95 per cent confidence level ULs for 0.16' disc and point-like hypotheses are shown with dashed green and light green lines, respectively. MAGIC 95 per cent confidence level ULs are displayed for both a point-like hypothesis (light orange) and a 0.16' radial extension (orange). The HAWC spectrum (dark blue) is obtained for the parameters given in Table 1. The light blue band indicates the HAWC spectrum taking into account 1σ systematic errors of 0.2 and 50 per cent in the photon index and flux, respectively.
our model, such a PWN would require an energy injection greater than $\sim 6.0 \times 10^{50}$ erg, which is already higher than $\dot{W}_p$. The low $B$ and $\dot{u}_B$ used in this calculation of $\dot{L}_{\text{cool}}$ provide maximum values for both $\dot{L}_{\text{cool}}$ and the injected pulsar energy, $\dot{W}_p$. Higher $\dot{u}_B$ value would produce higher losses and therefore, a smaller $\dot{W}_p$. However, an extremely high value for $\dot{u}_B$ (well above 2 eV cm$^{-3}$, the IR energy density observed around Cassiopea A and one of the highest for a Galactic TeV source) would be needed to decrease $\dot{W}_p$ below 10$^{50}$ erg. We do not consider more complex scenarios in which $E$ or $B$ change with time. The same parent population study was applied to 2HWC J1907+084 and we reach the same conclusions that corroborated the non-relation with the surrounding pulsars.

MAGIC and LAT ULs also help to constrain our understanding of the spectrum and morphology of these HAWC sources. The SEDs for the three candidate PWNe are shown in Fig. 3. MAGIC and Fermi-LAT analyses are computed with the photon index provided by HAWC (see Table 1). In the cases of 2HWC J2006+341 and 2HWC J1907+084, the MAGIC and LAT extended ULs are at the level of the HAWC spectrum considering HAWC systematic errors of 0.2 in the photon index and 50 per cent in the flux normalization. However, point-like hypotheses are in contradiction with HAWC results below energies of $\sim 4$ TeV and $\sim 900$ GeV, respectively. Therefore, it is expected that these two 2HWC sources are extended, with at least a radius of $\sim 0.16'$. On the other hand, both MAGIC and Fermi-LAT results on 2HWC J1852+013 are incompatible with the HAWC spectrum below energies of $\sim 10$ TeV.

These results can be understood in two ways: 2HWC J1852+013 is much more extended than the assumed radius of 0.16', which would increase MAGIC and Fermi-LAT ULs above the flux estimated by HAWC; or the source does not emit in the sub-TeV regime, consistent with the constraining ULs obtained by both MAGIC and the LAT. In the latter case, the spectral shaped of 2HWC J1852+013 would have a harder spectrum in the sub-TeV regime, and a minimum energy of around 10 TeV, instead of 400 GeV, should be assumed (see Table 1). To constrain the former case, we calculated LAT ULs for discs of larger radii. For a disc of 1.0' radius the LAT UL at energies $\geq 0.2$ TeV is within 1σ statistical errors of the measured HAWC flux, extrapolated to lower energies. However, this would also require a harder spectrum in the GeV regime so as to not exceed LAT ULs at lower energies. Additionally, as reported in the 2HWC catalogue Abeysekara et al. (2017c), there may be a significant contribution from diffuse Galactic emission at the location of 2HWC J1852+013 to which HAWC would be sensitive and MAGIC would not.

5 CONCLUSION

After the release of the 2HWC catalogue, MAGIC and Fermi-LAT performed dedicated analyses on three new TeV sources detected by the wide FoV observatory HAWC. None of them were detected at lower energies and no hotspot was found near them. However, owing to the increased time and good-quality data of most of the MAGIC and the Fermi-LAT observations, constraints on the extension of the sources were possible. With this aim, we performed both point-like and extended source searches. For 2HWC J2006+341 and 2HWC J1907+084, a radius of $\sim 0.16'$ is viable, given limits from the extended source search by MAGIC. For 2HWC J1852+013, MAGIC and Fermi-LAT results with respect to HAWC spectra suggest a much larger extension or a harder spectrum below $\sim 10$ TeV. Moreover, we find that none of the known pulsars in the vicinity of 2HWC J2006+341, 2HWC J1907+084 or 2HWC J1852+013 are likely to directly power these objects. It may be that these 2HWC sources are PWN created by as yet un-detected pulsars, or have some other origin such as a Galactic supernova remnant. More sensitive observations in the near future will provide valuable information on the nature of these sources and help to disentangle features in the crowded regions.

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