Discovery of a Supercluster at $z \sim 0.91$ and Testing the $\Lambda$CDM Cosmological Model

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ABSTRACT

The $\Lambda$CDM cosmological model successfully reproduces many aspects of the galaxy and structure formation of the universe. However, the growth of large-scale structures (LSSs) in the early universe is not well tested yet with observational data. Here, we have utilized wide and deep optical–near-infrared data in order to search for distant galaxy clusters and superclusters ($0.8 < z < 1.2$). From the spectroscopic observation with the Inamori Magellan Areal Camera and Spectrograph (IMACS) on the Magellan telescope, three massive clusters at $z \sim 0.91$ are confirmed in the SSA22 field. Interestingly, all of them have similar redshifts within $\Delta z \sim 0.01$ with velocity dispersions ranging from 470 to 1300 km s$^{-1}$. Moreover, as the maximum separation is $\sim 15$ Mpc, they compose a supercluster at $z \sim 0.91$, meaning that this is one of the most massive superclusters at this redshift to date. The galaxy density map implies that the confirmed clusters are embedded in a larger structure stretching over $\sim 100$ Mpc. $\Lambda$CDM models predict about one supercluster like this in our surveyed volume, consistent with our finding so far. However, there are more supercluster candidates in this field, suggesting that additional studies are required to determine if the $\Lambda$CDM cosmological model can successfully reproduce the LSSs at high redshift.

1. INTRODUCTION

Under the currently popular $\Lambda$CDM cosmology [Im et al. 1997; Riess et al. 1998; Perlmutter et al. 1999; Eisenstein et al. 2005], the large-scale structure (LSS) of galaxies emerges when the initial density fluctuations grow with time through gravitational attraction between galaxies. The $\Lambda$CDM cosmological model has been successful in reproducing the LSS at $z \sim 0$, showing the promise of the $\Lambda$CDM cosmology to explain our universe (e.g., Bahcall et al. 2003; Williamson et al. 2011; Benson et al. 2013).

However, the growth of LSSs has not been thoroughly tested yet with observational data at $z \gg 0$. LSSs stretch from several tens of Mpc to a few hundred Mpc, but there is a lack of data sets that are deep and wide enough to cover
such structures at high redshifts. So far, high-redshift LSS studies have been limited mostly to galaxy-cluster-scale structures ($\lesssim$1–2 Mpc), with mixed results. Some results show agreements with the $Λ$CDM cosmology models (Williamson et al. 2011; Bayliss et al. 2014), but others suggest too many massive clusters at $z \gtrsim 1$ (Jee et al. 2009; Gonzalez et al. 2012; Kang & Im 2013).

With the advance of large and deep imaging surveys, it is now possible to extend the test of the cosmological formation of LSSs to scales much larger than before. Specific predictions have been made about superclusters at high redshift. A supercluster represents the most massive structure in the universe with sizes of up to $\sim$100–200 Mpc, containing filaments, multiple galaxy clusters and groups. Several studies have pointed out that superclusters are useful objects to test cosmological predictions. Specific predictions have been made about superclusters at high redshift. A supercluster at $z \gtrsim 0.9$ has 8 member clusters and groups confirmed spectroscopically (Rosati et al. 2016, in preparation). The IMS is a deep (J $\approx$ 25 AB mag) and wide ($\sim$120 deg$^2$) near-infrared (NIR) imaging survey that combines deep J-band imaging data with other optical/NIR survey data, making it possible to find LSSs at $z \sim 1$. Here, we report the discovery of a new, massive supercluster at $z \sim 0.9$ in the SSA22 area as the first result, and discuss if such an LSS is compatible with cosmological simulation predictions.

We adopt cosmological parameters for the flat universe with $\Omega_m = 0.27$, $H_0 = 71$ km $s^{-1}$ Mpc$^{-1}$, and $σ_8 = 0.8$. All magnitudes are in the AB system. In addition, all distance scales are physical scales based on the angular diameter distance, unless otherwise noted.

2. DATA AND CLUSTER FINDING

2.1. Photometric Catalog

Our work is based on wide and deep data sets for the SSA22 field ($α = 22^h 17^m 00^s$ and $δ = 00^\circ 20' 00''$) from the Canada-France-Hawaii Telescope (CFHT) Legacy Survey (CFHTLS), the United Kingdom Infrared Telescope (UKIRT) Infrared Deep Sky Survey (UKIDSS, Lawrence et al. 2007), Deep eXtragalactic Survey (DXS: Edge et al. 2016, in preparation), and the IMS. Although UKIDSS DXS ($J$- and $K$-bands) and IMS mapped nearly the entire CFHTLS–W4 (ugriz-bands) area (25 deg$^2$), the effective area is $\sim$20 deg$^2$ after excluding regions such as halos and spikes of bright stars. The 80% point-source detection limits are $u^* \sim 25.2$, $g' \sim 25.6$, $r' \sim 25.0$, $i' \sim 24.9$, and $z' \sim 23.9$ for the CFHTLS-3 and $J \sim 23.7$ and $K \sim 23.2$ for UKIDSS DXS and IMS (Kim et al. 2011, 2013). Sources were detected using SExtractor (Bertin & Arnouts 1996) in dual mode, and the unconvolved J-band images were used for the detection and to measure the J-band total magnitude. In addition, 2$''$ diameter apertures were applied to PSF matched images to derive aperture magnitudes for the color measurement. For this work, we applied a magnitude cut of $J = 23.2$ that is the 90% point-source completeness limit of the J-band data. Photometric redshifts ($z_{phot}$) were derived using the Le Phare software (Arnouts et al. 1999; Ilbert et al. 2006) after training the data using spectroscopic redshifts with flags of 3, 4, 23, and 24 from the VIMOS VLT Deep Survey (VVDS: Le Fèvre et al. 2005; Garilli et al. 2008). A measured redshift accuracy ($σ_{pz}$) by the normalized median absolute deviation is $Δz/(1 + z) = 0.038$, and the outlier fraction is $<5\%$. Details of the procedures are described in Kim et al. (2013).

1http://www.cfht.hawaii.edu/Science/CFHTLS/

2http://terapix.iap.fr/
2.2. Finding MSGs

Using all objects in the photometric redshift catalog, we searched for massive structures of galaxies (MSGs, galaxy clusters and groups) between $z = 0.8$ and $z = 1.2$. After splitting galaxies into redshift bins from $z_{bin} = 0.8$ to $z_{bin} = 1.2$ with an increment of 0.02 and a bin size of $|z_{phot} - z_{bin}| < \sigma_{pz}(1 + z_{bin})$ based on the best-fit photometric redshift, the Voronoi Tessellation technique (Ebeling & Wiedenmann 1993; Soares-Santos et al. 2011) is applied to measure a local density ($\rho = 1/\text{area}_{cell}$) for each galaxy, which is converted into the normalized cell density, $\delta = \rho/\rho_{\text{median}}$. Galaxies are identified to be in an overdense region if its $\delta$ value is above a threshold, $\delta_{\text{thres}}$, following the prescription described in Soares-Santos et al. (2011). The $\delta_{\text{thres}}$ values are determined for each redshift bin, and they are found to vary between $\delta = 1.9$ and 2.0. The threshold corresponds to approximately $4\sigma$ above the mean density, if we fit the $\delta$ distribution with a Gaussian function in linear scale at $\delta < 1^3$ We group galaxies whose cells have $\delta > \delta_{\text{thres}}$ and are adjacent to each other as a possible overdense area. Then, we classify the overdense area as an MSG candidate if it has a probability over 95% that the signal is not due to random fluctuations (Eq. (3) of Soares-Santos et al. 2011). The sky position of a MSG candidate is assigned as the coordinate of the galaxy with the highest density. Also, the redshift is taken to be the median redshift of its galaxies within 1 Mpc from the MSG candidate center. MSG candidates from different redshift bins are merged into a single candidate if the projected separation between them is less than 2 Mpc, and their redshift bins overlap with each other.

We estimated the fraction of bona fide member galaxies in this approach using a galaxy mock catalog of GALFORM (Cole et al. 2000; Merson et al. 2013). For this, we randomly scatter mock galaxy redshifts with $\sigma_{pz}$ above, and then select galaxies with $|z_{phot} - z_{cen}| < \sigma_{pz}(1 + z_{cen})$ and within 1.0 and 1.5 Mpc radii from central galaxies of 346 massive halos at 0.85 < $z_{cen}$ < 1.15, where $z_{cen}$ indicates the halo redshift. The average fractions of bona fide members among selected galaxies are 50% and 30% for 1.0 and 1.5 Mpc, respectively. Therefore, we select MSG candidates, only if they have at least 25 galaxies ($N_{1.5\text{Mpc}} \geq 25$) within a 1.5 Mpc radius and the photometric redshift uncertainty and with $J < J^* + 1$ (where $J^*$ is characteristic magnitude). In total, there are 691 MSG candidates.

In order to identify supercluster candidates, we count the number of MSG candidates within a 10 Mpc radius and the photometric redshift uncertainty from each MSG candidate. Then, supercluster candidates are chosen as a group of at least 10 MSG candidates. Through this process, we find two supercluster candidates at the median photometric redshift of $\sim 0.89$ and another at $\sim 0.92$.

3. CLUSTER CONFIRMATION

3.1. IMACS Observation and Redshift Determination

Multi-object spectroscopy was performed on 2014 September 23, using the Inamori Magellan Areal Camera and Spectrograph (IMACS) on the Magellan/Baade telescope in its $f/2$ mode of a field that covers a 27′.4 diameter field of view at $\alpha = 22^h 13^m 08^s$ and $\delta = 00^d 40^m 24^s$. Of supercluster candidates described in the previous section, we chose the target field due to its unusually high concentration in a small area: six prominent ($N_{1.5\text{Mpc}} > 40$) and nine less significant ($N_{1.5\text{Mpc}} < 40$) MSG candidates, within a photometric redshift range of $0.85 < z_{phot} < 0.96$. Slitlets were assigned to galaxies in prominent candidates first, and then to those in less significant candidates. In order to choose target galaxies for the spectroscopy, we used the probability distribution function (PDF) of photometric redshifts from the Le Phare software. The integrals of the normalized PDFs within the uncertainty range ($\sigma_{pz}$) from the candidate redshift were calculated as the probability for each galaxy belonging to the cluster (Brunner & Lubin 2006; Papovich et al. 2010; Brodwin et al. 2013). Galaxies with probabilities >0.5 were selected as potential members. Spectra of potential members were taken using the 200 lines mm$^{-1}$ grism with the WB5600–9200 filter ranging from 5600A to 9200A. One slit mask was used for the observation with 1′′×6′ slitlets. In total, 320 slitlets were assigned for galaxies including potential MSG members (80%) and field galaxies.
(20%). The spectral resolution was $\lambda/\Delta \lambda \sim 600$. The total on-source integration time was 2.5 hours (30 minutes x 5) under $\sim 0''$.9 seeing.

We used the Carnegie Observatories System for MultiObject Spectroscopy (COSMOS) to reduce the IMACS spectroscopic data. The procedure includes standard reduction algorithm, wavelength calibration, and sky subtraction. We extracted one-dimensional spectra for each source from two-dimensional spectra stacked by the COSMOS pipeline. The flux calibration was performed using an F5-type star that was also included in the slit mask.

The redshift of each galaxy was determined with the SpecPro software [Masters & Capak 2011]. We mainly used the emission and absorption lines of [OII]3727, Ca H&K, the 4000 Å break, the G-band, and the Balmer lines (H$\delta$ and H$\gamma$) for this. If only a single emission line was detected, we considered this as the [OII] line. If the identified line was not [OII], e.g., such as H$\delta$, H$\beta$, [OIII] or H$\alpha$, it would likely be accompanied by another line at shorter or longer wavelengths and be at a redshift that is difficult to explain the continuum shape (i.e., at a redshift that is very different from photometric redshifts). We successfully determined the redshift of 217 galaxies, implying a success rate of $\sim$70%. The success rate is $\sim$80% for galaxies at $i_{AB} \leq 22.5$. Among successful spectroscopic measurements, 51% comes from a single emission, and 7% comes from Ca H&K absorption lines.

### 3.2. Discovery of Supercluster at $z = 0.91$

Using the galaxies with spectroscopic redshifts, we determine the membership of each galaxy. For this process, we follow the iterative algorithm described in [Lubin et al. 2002]. First, we select galaxies within a 1 Mpc radius from the cluster position determined by the photometric redshift method in § 2.2. Then, we calculate the bi-weight mean ($z_{cl}$) and scale ($\sigma_z$) of redshifts of these galaxies [Beers et al. 1990]. We exclude galaxies with $|z - z_{cl}| > 3\sigma_z$ or the relative rest-frame radial velocity greater than 3500 km s$^{-1}$. This process is repeated until no more galaxies are excluded. Finally, the dispersion ($\sigma_z'$) is calculated by the gapper method due to the small number of members and then converted into the velocity dispersion ($\sigma_v$). Uncertainties in $\sigma_v$ are estimated by the Jackknife resampling. Through this process, we identify three galaxy clusters at $z \sim 0.91$ with $\sigma_v > 470$ km s$^{-1}$, which corresponds to $M_{200} > 1.1 \times 10^{14} M_{\odot}$ at the cluster redshift [Carlberg et al. 1997; Demarco et al. 2010]. Table 1 lists the properties of the confirmed clusters based on two different radii of 1 Mpc and 1.5 Mpc for the comparison. Independently, we estimated the uncertainty in $\sigma_v$ by randomly selecting 100 times 11 and 7 spectroscopic members within a 1 Mpc radius of the centers of two galaxy clusters at $z \sim 1.2$ containing $\sim$20-30 galaxies each with $\sigma_v = 490-650$ km s$^{-1}$ [Muzzin et al. 2009]. We find that the standard deviation from this exercise to be 160 km s$^{-1}$ and 320 km s$^{-1}$ for the 11 and 7 member cases, respectively, which are consistent with or smaller than the Jackknife resampling errors in Table 1.

Figure 1a shows example spectra of cluster members in the three confirmed clusters. Vertical lines indicate locations of [OII] (blue), Ca H&K (red), and H$\delta$ (magenta) lines. Figure 1b displays the spectroscopic redshift distribution of confirmed members for each cluster. The fraction of members of which spectroscopic redshifts are determined by a single line is 55%, 29%, and 0% for IMSCl J2212+0045, IMSCl J2213+0052 and IMSCl J2213+0048, respectively. Note that the velocity dispersion based only on red galaxies could decrease as much as $\sim$50% of that from both blue and red galaxies [Gal et al. 2008]. When we did a similar analysis on two clusters where at least 5 and 4 red galaxies ($(r-i) > 0.9$ and $(i-z) > 0.5$) are available (J2212+0045 and J2213+0048), $\sigma_v$ becomes 200±149 and 1117±464 km s$^{-1}$, respectively. This suggests that the $\sigma_v$ values in Table 1 could be overestimated. Figure 2 shows pseudo color images for the confirmed clusters with photometric (circles) and spectroscopic (squares) members.

For the SSA22 field, [Durret et al. 2011] also identified candidate galaxy clusters using photometric redshifts derived from the optical CFHTLS data. Of our clusters, IMSCl J2213+0045 and IMSCl J2213+0048 have counterparts in their list, considering a matching radius of 3' ($\sim$1.5 Mpc at $z = 0.91$) and $\Delta z < 0.1$.

Interestingly, all the new clusters are massive
and located at redshifts of $z \sim 0.908$–0.920. In addition, the maximum projected angular separation between galaxy clusters is 16.78, which corresponds to $\sim 15$ Mpc (comoving) at the cluster redshift. The proximity of the three massive clusters suggests that this is a supercluster. Compared to galaxy clusters in the Cl 1604 supercluster with a maximum velocity dispersion of $\sim 800$ km s$^{-1}$, the supercluster presented here may encompass one or more clusters that are more massive and is more comparable to the RCS 2319+00 supercluster members (710–1200 km s$^{-1}$) and members in the Elais-N1 field (660–1000 km s$^{-1}$). Additionally, the maximum projected angular separation is similar to that for the three most massive clusters in Cl 1604 (16.24), larger than RCS 2319+00 (7.62), and smaller than the Elais-N1 supercluster (32.45). This new supercluster in the SSA22 field may be one of the most massive structures ever found at $z \sim 0.9$. We summarize the properties of the known superclusters in Table I.

4. DISCUSSION

4.1. Large-scale Structure

In order to see if there are more structures beyond the confirmed clusters through a galaxy density map, we apply the Voronoi tessellation technique described in § 2.2 to the redshift bin at $z = 0.914$, the mean redshift of the confirmed clusters. Then, we make a grid map with a grid cell size of 500 kpc across the entire SSA22 area and calculate the mean value of local densities ($\delta$ in § 2.2) for each grid.

Figure 3 shows the map of local densities at $z \sim 0.914$ around the confirmed clusters over a 3.5 deg $\times$ 2.9 deg area ($\sim 190$ Mpc $\times$ 160 Mpc, comoving). Small and large red circles indicate the positions of the newly confirmed clusters and the IMACS pointing, respectively. Intriguingly, the density contours extend toward the south and northeast with the structure spanning from (R.A., decl.)=(333.12, -0.1) to (334.15, 1.5). The projected angular size of this structure is $\sim 114.2$ corresponding to $\sim 54$ Mpc or $\sim 103$ Mpc (comoving) at this redshift. We also overlay galaxies (orange points) with spectroscopic redshifts of $0.90 < z_{\text{spec}} < 0.92$ from the VVDS and the VIMOS Public Extragalactic Redshift Survey that are deemed reliable (flags 2–9 and 22–29; VIPERS; Garilli et al. 2014 Guzzo et al. 2014). Comparing the density map and the distribution of spectroscopic samples, it seems that the supercluster extends the northeast direction. Note that no spectroscopic redshifts are available in the southwest area. The density map suggests that the supercluster possibly extends to a much larger scale.

4.2. Comparison with Models

Here, we examine if the existence of the new supercluster at $z = 0.91$ can be explained with ΛCDM models.

First, we search for dark matter halos grouped similarly to our confirmed clusters from the Millennium simulation (Springel et al. 2005) with the WMAP-7 cosmology (Guo et al. 2013). We use 15 snapshots from $z = 1.77$ to $z = 0.51$, each with a 0.32 Gpc$^3$ cube (i.e., five times the volume of our data at $0.8 < z < 1.2$). At $z = 1.08$, the first structures with properties comparable to the new supercluster form, i.e., those containing at least three halos, each in excess of $M_{200} > 1.1 \times 10^{14} M_\odot$. Two such structures appear at this redshift. By $z = 0.8$, three or four such structures have formed. When translated to our survey volume, the simulation suggests $\sim 0.6$–0.8 superclusters at $0.8 < z < 1.2$. If we use the Millennium simulation with $\sigma_8 = 0.9$, the number is comparable in this redshift range.

Second, we calculate the predicted number of superclusters based on the supercluster mass function of Lim & Lee (2014). As a conservative estimate, we set the supercluster mass at $10^{15} M_\odot$. Using their mass function at $z = 1$, the predicted number of superclusters with $> 10^{15} M_\odot$ in the 20 deg$^2$ area with $0.8 < z < 1.2$ is $\sim 0.6$.

The expected numbers of superclusters from the models are consistent with the number of superclusters we identified so far. However, there are still two more supercluster candidates in the SSA22 field, and extended structures as discussed in § 4.1 may contain more superclusters. On the other hand, considering the uncertainty of cluster masses, the clusters can be lighter. If so, the observation and the model predictions can be reconciled since lighter superclusters are more abundant than heavier ones in models. To understand if there is any tension between observed superclusters and ΛCDM models, it is necessary to do a more thor-
ough analysis of larger cosmological simulations and an intensive spectroscopic mapping of these large structures.

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REFERENCES
Arnouts, S., Cristiani, S., Moscardini, L., et al. 1999, MNRAS, 310, 540
Ascaso, B., Lemaux, B. C., Lubin, L. M., Gal, R. R., Kocevski, D. D., Rumbaugh, N., & Squires, G. 2014, MNRAS, 442, 589
Bahcall, N. A., Dong, F., Hao, L., Bode, P., Annis, J., Gunn, J. E., & Schneider, D. P. 2003, ApJ, 599, 814
Bayliss, M. B., Ashby, M. L. N., Ruel, J., et al. 2014, ApJ, 794, 12
Beers, T. C., Flynn, K., & Gebhardt, K. 1990, AJ, 100, 32
Benson, B. A., de Haan, T., Dudley, J. P., et al. 2013, ApJ, 763, 147
Bertin, E., & Arnouts S. 1996, A&AS, 117, 393
Brodwin, M., Stanford, S. A., Gonzalez, A. H., et al. 2013, ApJ, 779, 138
Brunner, R. J., & Lubin, L. M. 2000, AJ, 120, 2851
Carlberg, R. G., Yee, H. K. C., & Ellingson, E. 1997, ApJ, 478, 462
Cole, S., Lacey, C. G., Baugh, C. M., & Frenk, C. S. 2000, MNRAS, 319, 168
Demarco, R., Wilson, G., Muzzin, A., et al. 2010, ApJ, 711, 1185
Durret, F., Adami, C., Cappi, A., et al. 2011, A&A, 535, 65
Ebeling, H., & Wiedenmann, G. 1993, Phys. Rev. E, 47, 704
Einasto, M., Liivamägi, L. J., Tago, E., Saar, E., Tempel, E., Einasto, J., MartínezV. J., & Heinämäki, P. 2011, A&A, 532, 5
Eisenstein, D. J., Zehavi, I., Hogg, D. W., et al. 2005, ApJ, 633, 560
Faloon, A. J., Webb, T. M. A., Ellinson, E., et al. 2013, ApJ, 768, 104
Gal, R. R., Lemaux, B. C., Lubin, L. M., Kocevski, D., & Squires, G. K. 2008, ApJ, 684, 933
Garilli, B., Guzzo, L., Scudellaro, M., et al. 2014, A&A, 562, 23
Garilli, B., Le Fèvre, O., Guzzo, L., et al. 2008, A&A, 486, 683
Gilbank, D. G., Yee, H. K. C., Ellingson, E., Hicks, A. K., Gladders, M. D., Barrientos, L. F., & Keeney, B. 2008, ApJ, 677, L89
Gonzalez, A. H., Stanford, S. A., Brodwin, M., et al. 2012, ApJ, 753, 163
Guo, Q., White, S., Boylan-Kolchin, M., De Lucia, G., Kauffmann, G., Lemoine, G., Li, C., Springel, V., & Weinmann, S. 2013, MNRAS, 428, 1351
Guzzo, L., Scodeggio, M., Garilli, B. R., et al. 2014, A&A, 566, 108
Ilbert, O., Arnouts, S., McCracken, H. J., et al. 2006, A&A, 457, 841
Im, M., Griffiths, R. E., & Ratnatunga, K. U. 1997, ApJ, 475, 457
Jee, M. J., Rosati, P., Ford, H. C., et al. 2009, ApJ, 704, 672
Kang, E., & Im, M. 2015, JKAS, 48, 21
Kim, J. -W., Edge, A. C., Wake, D. A., & Stott, J. P. 2011, MNRAS, 410, 241
Kim, J. -W., Im, M., Lee, S. -K., et al. 2015, ApJ, 806, 189
Lawrence, A., Warren, S. J., Almaini, O., et al. 2007, MNRAS, 379, 1599
Le Fèvre, O., Vettolani, G., Garilli, B., et al. 2005, A&A, 439, 845
Lemaux, B. C., Gal, R. R., Lubin, L. M., et al. 2012, ApJ, 745, 106
Lim, S., & Lee, J. 2014, ApJ, 783, 39
Merson, A. I., Baugh, C. M., Helly, J. C., et al. 2013, MNRAS, 429, 556
Papovich, C., Momcheva, I., Willmer, C. N. A., et al. 2010, ApJ, 716, 1503
Perlmutter, S., Aldering, G., Goldhaber, G., et al. 1999, ApJ, 517, 565
Riess, A. G., Filippenko, A. V., Challis, P., et al. 1998, AJ, 116, 1009
Rosati, P., Stanford, S. A., Eisenhardt, P. R., Elston, R., Spinrad, H., Stern, D., & Dey, A. 1999, AJ, 118, 76
Soares-Santos, M., de Carvalho, R. R., Annis, J., et al. 2011, ApJ, 727, 45
Springel, V., White, S. D. M., Jenkins, A., et al. 2005, Nature, 435, 629
Swinbank, A. M., Edge, A. C., Smail, I., et al. 2007, MNRAS, 379, 1343
Williamson, R., Benson, B. A., High, F. W., et al. 2011, ApJ, 738, 139
Wray, J. J., Bahcall, N. A., Bode, P., Boettiger, C., & Hopkins, P. F. 2006, ApJ, 652, 907
Wu, P. -F., Gal, R. R., Lemaux, B. C., Kocevski, D. D., Lubin, L. M., Rumbaugh, N., & Squires, G. K. 2014, ApJ, 792, 16
Fig. 1.— (a) Example IMACS spectra of the confirmed cluster members. The cluster IDs and the spectroscopic redshifts are noted in each panel. The vertical lines mark the [OII]3727, Ca H&K and Hδ4102 lines at the noted redshift. (b) Redshift distribution of spectroscopic members within 1 Mpc. The dotted lines show Gaussian distributions based on $z_{cl}$ and $\sigma'_z$ (see § 3.2 for details).

Fig. 2.— Pseudo-color images ($gKL$-bands) of the confirmed clusters. These images were arbitrarily scaled for display purposes. The field of view is 5\arcmin$\times$5\arcmin corresponding to 2.4 Mpc$\times$2.4 Mpc at the cluster redshift. The circles and the squares are for photometric and spectroscopic members, respectively. Yellow curves show contours for $< \delta >$=2, 4, 6, 7, and 10 at $z \sim 0.91$ (§ 4.1).
Table 1: Summary of confirmed clusters in this study with applying two different radii (top) and previously reported superclusters at $z \sim 0.9$ (bottom).

| Cluster       | R.A. (J2000) | decl. (J2000) | Radius (Mpc) | $n_{\text{cl}}$ | $n_{\text{member}}$ | $z_{\text{cl}}$ | $\sigma_v$ (km s$^{-1}$) | $M_{200} \left( \times 10^{14} M_\odot \right)$ |
|---------------|-------------|---------------|--------------|-----------------|---------------------|-----------------|--------------------------|-----------------------------------------------|
| IMSci J2212+0045 | 22:12:28    | 00:45:06      | 1.0          | 19              | 11                  | 0.9170±0.0008    | 474±152                 | 1.1$^{+0.8}_{-0.6}$                           |
| IMSci J2213+0052 | 22:13:02    | 00:52:02      | 1.0          | 13              | 7                   | 0.9166±0.0026    | 884±160                  | 7.2$^{+3.3}_{-2.5}$                           |
| IMSci J2213+0048 | 22:13:31    | 00:48:42      | 1.0          | 17              | 7                   | 0.9085±0.0041    | 1298±310                 | 23.0$^{+20.7}_{-12.8}$                        |

| Supercluster   | R.A.$^a$ | decl.$^a$ | $z$ | $n_{\text{cluster}}$ | $\sigma_v$$^b$ | Size$_{\text{projected}}$$^b$ | $<n_{\text{galaxy}}>_6$ | Ref. |
|----------------|---------|---------|-----|---------------------|--------------|-----------------------------|-------------------------|------|
| Cl 1604        | 16:04:23| 43:13:08| 0.85-0.94 | 8       | 688-818            | 17$^f$         | ∼70                      | Wu+14$^c$           |
| Swinbank+07    | 16:08:27| 54:35:47| 0.89 | 5                   | 730-1030     | 32$^f$          | ∼11                    | Swinbank+07         |
| RCS 2319+00    | 23:19:53| 00:38:04| 0.90 | 8                   | 714-1202     | 8$^f$           | ∼16                    | Falloon+13          |

$^a$Coordinate for the most massive cluster.
$^b$Values based on the three most massive clusters.
$^c$Member galaxies within two times the virial radius in [Wu et al. 2014] and 1 $h^{-1}$Mpc in [Gal et al. 2008].

Fig. 3.— Overdensity contour at $z = 0.914$ around the confirmed clusters. Small and large red circles show the confirmed clusters and the IMACS field of view, respectively. Dashed and dotted-dashed boxes indicate survey boundaries of VVDS and VIPERS, respectively. The orange points are galaxies with spectroscopic redshifts between $z = 0.90$ and 0.92 from VVDS and VIPERS. More MSG candidates exist at northeast of the supercluster, and spectroscopic samples from VVDS and VIPERS seem to connect the supercluster and the MSG candidates.