The surface brightness of dark matter: 
unique signatures of neutralino annihilation in the Galactic halo.

Carlos Calcáneo–Roldán* & Ben Moore†.

Department of Physics, Durham University, Science laboratories, Durham, DH1 3LE, UK.

(October 29, 2018)

We use high resolution numerical simulations of the formation of cold dark matter halos to simulate the background of decay products from neutralino annihilation, such as gamma-rays or neutrinos. Halos are non-spherical, have steep singular density profiles and contain many thousands of surviving dark matter substructure clumps. This leads to several unique signatures in the gamma-ray background that may be confirmed or rejected by the next generation of gamma-ray experiments. Most importantly, the diffuse background is enhanced by over two orders of magnitude due to annihilation within substructure halos. The largest dark substructures are easily visibly above the background and may account for the unidentified EGRET sources. A deep strip survey of the gamma-ray background would allow the shape of the Galactic halo to be quantified.

I. INTRODUCTION

Determining the nature of dark matter is of fundamental importance to both Astronomy and Particle Physics. Both theory and observational data currently favour a universe with a matter density that is dominated by non-baryonic particles. Many candidates have been proposed: some are known to exist, others are more speculative (e.g., Ref. 1 and references therein). Structure formation in a universe dominated by cold dark matter (CDM) has been extensively tested against observations and the model has proven highly successful at reproducing the large scale properties and distributions of galaxies 2,3. On the non-linear scales of galactic halos it remains to be confirmed whether the model can successfully reproduce the observational data 1,8.

Direct detection in the laboratory is the ultimate technique for verifying the existence of dark matter particles (see, e.g., Ref. 1 and references therein). However, even the most popular candidate for dark matter, the neutralino, has a cross section that spans many orders of magnitude and the current laboratory searches are only just becoming sensitive to the cosmologically interesting parameter range. Presently, Astronomical observations provide the best insights into the nature of the dark matter, furthermore direct detection relies on the existence of a smooth component of dark matter.

Within the next few years indirect detection of neutralinos will provide interesting constraints on their possible cross-section and masses. Neutralino-neutralino annihilation produces observable photons (as well as a host of other particles) that may be observed as a diffuse gamma-ray background from the halo surrounding the Milky Way as discussed in Refs. 1,15, and more recently, in Refs. 14,19. Renewed interest in these predictions has recently arisen because of an unexplained component of diffuse high energy photons in the Egret data (e.g. Ref. 20), and also the possibility of an excess from the center of the Galaxy itself 21, unexpected clumpy emission and the unresolved "discrete sources" 22. Progress in this area will result from several new and sensitive gamma-ray surveys such as GLAST 23 and VERITAS 24.

The efficiency of the annihilation process is strongly dependent on the local density and the cross-section of the neutralino. Many authors have calculated the expected flux from the Galactic halo using simple models for the expected mass distribution of neutralinos within the Galaxy (see, e.g., Refs. 14,19) or from its satellites 14,19.

Advances in computational Cosmology have lead to several recent breakthroughs that have direct relevance to the detection of dark matter. In particular, the numerical resolution that can be achieved using parallel computational techniques is now sufficient to study the internal structure of dark matter halos that form within a cosmological context. The results of these simulations have important implications for indirect (and direct) detection of dark matter candidates. Most significantly for particle-particle annihilation, we are now confident that the central density profile of CDM halos follows a singular power law down to small scales 25,26. Thus we may expect a point like source of mono-chromatic gamma-rays emanating from the center of the Milky Way, where the annihilation rate will be very high.

A second fundamental prediction of the CDM model is that previous generations of the merging hierarchy survive within halos 25. Halos that accrete into larger systems may be tidally stripped of most of their mass, however their dense central cores survive and continue to orbit within the parent halos. This may present some problems for the CDM model since the predicted number of satellites within the Milky Way’s halo is 50–100 times as many as observed 3. If the CDM model is correct,
then only a fraction of these satellites must have formed stars and most of the substructure remains as dark objects within the Galactic halo.

The possibility of an enhanced gamma-ray background from dark matter substructure was explored by Bergström et al. [29], who made simple assumptions as to the mean density and abundance of such clumps. We can now use the high resolution N-body simulations to directly measure these quantities. The simulations also allow us to study the influence of the halo shapes on the diffuse gamma-ray background and the intensity of the central halo emission that arises from the singular dark matter density profiles. This paper is organized as follows. In Section II we explore the gamma-ray background that results from the smoothly distributed component of dark matter using both analytical and simulated halos. In Section III we focus on the substructure within halos. Our conclusions are summarised in Section IV.

II. THE SKY DISTRIBUTION OF THE GAMMA-RAY BACKGROUND

In what follows we will consider a flux of photons (or other particles) that are a by-product of the annihilation of dark matter particles within the smooth component of dark matter that surrounds the Galaxy. It is not our intention to discuss the details of neutralino interactions, a complete overview on these processes (and super-symmetric matter in general) can be found in Refs. [30,31].

A. Model neutralino halos

We calculate the gamma-ray flux along a given line of sight through a spherically symmetric galactic halo using:

\[ \phi(\psi) = \frac{K}{4\pi} \int \rho(l)^2 d\ell(\psi) \]  

(1)

where \( \psi \) is the angle between the direction of galactic center and observation; \( \rho \), the density of dark matter at distance \( l \) from the observer. We have summed up the dependence of the flux on neutralino mass and interaction cross section in the constant \( K \). This is enough scope for the present discussion - it is straightforward to take our results and input a neutralino cross-section, \( <\sigma v> \), and mass, \( M_\chi \) to determine the absolute gamma-ray flux (where \( K \) is defined to be \( <\sigma v>/M_\chi^2 \)). Our results can also be used to infer the sky distribution of other products of the annihilation, such as neutrinos or positrons.

The line of sight distance, \( l \), is related to the radial distance from the halo center, \( r \), via

\[ r^2 = l^2 + R_o^2 - 2lR_o \cos(\psi) \]

where \( R_o \) is our galacto-centric distance, taken here to have the IAU standard value of \( R_o = 8.5kpc \) [32], and \( \psi \) is related to galactic coordinates \((\ell, b)\) through

\[ \cos(\psi) = \cos(\ell) \cos(b). \]

For the halo density profile, \( \rho(r) \), we take the latest results from the highest resolution numerical simulations of galactic halos carried out to-date [3]. These authors simulated 6 different galactic mass halos with force resolution of 0.5 kpc and mass resolution of \( 10^5 M_\odot \). (Throughout the paper we will use the Hubble constant value of \( H_o = 100 h km/s/Mpc \) and \( h = 0.5 \); as adopted for the simulations.) The best fitting density profile to these data is (subscript moore):

\[ \rho_{\text{moore}}(r) = \frac{\rho_\text{moore}'}{(r/a)^{1.5}(1 + (r/a)^{1.5})}. \]

(2)

Where \( r \) is the distance from the halo center, \( a = r_{200}/r_{\text{moore}} \) the scale radius for halos of mass \( \approx 1 \times 10^{12} M_\odot \). The virial radius of our fiducial Galactic halo, \( r_{200} \approx 300 \) kpc, is defined as the radius of a sphere at which the mean overdensity is 200 times the cosmological mean density. (A central density profile of slope -1.5 on galactic scales was also found by Jing & Suto [33] and confirmed as an asymptotic slope by Ghigna et al. [34].)

We also compare this profile with that determined by Navarro et al. [26] using a sequence of lower resolution studies (subscript nfw) (the main difference being that the central dark matter density profile has a slope of -1):

\[ \rho_{\text{nfw}}(r) = \frac{\rho_\text{nfw}'}{(r/a)^2(1 + (r/a)^2)^2}. \]

(3)

and the modified isothermal profile with a constant density core (subscript is):

\[ \rho_{\text{is}}(r) = \frac{\rho_\text{is}'}{[1 + (r/a)^2]^{3/2}}. \]

(4)

The scale radius, \( a \), is determined directly from the numerical simulations, except for the modified isothermal model which we normalise to match the observational rotation curve data (as in Ref. [33]): \( a_{\text{is}} = 24.3 \) kpc, \( a_{\text{nfw}} = 27.7 \) kpc and \( a_{\text{moore}} = 33.2 \) kpc (this radius is directly related to the concentration parameter, \( c = r_{200}/a \)). We normalise each density profile such that the peak circular velocity, \( v_{\text{peak}} = 200 \) km/s (the maximum of the \( v_r = \sqrt{GM/r^2} \) curve), which gives:

\[ \rho_{\text{is}}' = 4.96 \times 10^9 M_\odot \text{ kpc}^{-3}, \rho_{\text{nfw}}' = 5.11 \times 10^9 M_\odot \text{ kpc}^{-3} \]

and \( \rho_{\text{moore}}' = 1.64 \times 10^9 M_\odot \text{ kpc}^{-3} \). We plot the effective circular velocity profiles and density profiles of these model halos in Fig. 3(a) and Fig. 3(b) respectively.

In Fig. 3 we plot the flux, \( \phi \), along the line of sight through a spherical Milky Way halo using the above density profiles as the observer looks towards the Galactic center at \( \psi = 0^\circ \), to the Galactic anticenter at \( \psi = 180^\circ \). As expected, the central annihilation flux...
depends strongly on the form of the inner density profile. At an angle of five degrees from the Galactic center, the ratio of fluxes from the three different profiles, moore:nfw:is are 1000:1:1.

The peak central value depends upon the distance from the Galactic center that we are willing to consider integrating from – the flux slowly diverges for the density profile in Eq. (2). However, within a given radius, most of the neutralinos would have self annihiliated leaving a tiny constant density core. We can estimate the size of this core using \((nσv)^{-1} = t_h\), where \(t_h \approx 10\) Gyrs is the Hubble time. Taking a typical cross section, \(σv = 10^{-30}\)cm\(^3\)s\(^{-1}\), and adopting the Moore et al. density profile we find that the annihilation radius within the Milky Way is approximately \(4 \times 10^{-7}\) parsecs \(\approx 10^{-12}r_{200}\).

The total flux that arises within 5 degrees of the Galactic center using the Moore et al. density profile is a factor of 20 larger than that found using the NFW profile (both integrated down to the annihilation radius calculated above).

**B. Comparison with high resolution CDM simulations**

We can use the numerical simulations to compare directly with the above predictions that were obtained assuming spherical symmetry. We refer the reader to Moore et al. (Ref. [3]) for details of the numerical simulations. To construct the expected gamma-ray sky maps we choose a simulated dark matter halo at a redshift \(z=0\) that has a peak circular velocity of \(\sim 200\text{kms}^{-1}\) and a total mass, within the virial radius, \(r_{200} = 300\text{kpc}\), of \(1 \times 10^{12}M_\odot\). This simulated halo is from the “Local Group” simulation and is close to our fiducial Milky Way cold dark matter halo that we adopted in the previous section.

N-body simulations attempt to simulate a collisionless fluid of dark matter using discrete massive particles. We calculate the local density at the position of each particle by averaging over its nearest 64 neighbours. The observer is placed 8.5 kpc from the halo center (defined using the most bound particle in the simulation) and we sum up the flux of annihilation products along each line of sight using the discrete equivalent to Eq. (1):

\[
\Phi(\ell, b) = \frac{K}{\Omega} \sum_{LOS} n\rho^2(\ell, b)\Delta r_i(\ell, b)
\]

where \(\ell, b\) are galactic longitude and latitude respectively. The flux is binned in angular windows of size \(\Omega = 1° \times 1°\) and in the radial direction in fixed increments \(\Delta r_i = 1\text{kpc}\).

The simulated dark matter halos are typically flattened oblate or prolate systems [16]. We do not know a-priori in which axis the stellar disk would be located, therefore we show two all-sky maps using the same dark matter halo but viewed using two different locations for the “observer”: Figure 3(a) and Fig. 3(b) has the observer located on the short and long axes respectively. Both of these plots show the enhanced brightening towards the halo center, as well as some clumpy substructure in the halo itself. Note that both the central halo and the centers of the substructure halos are artificially “dimmed” in these plots due to the numerical resolution \(\sim 0.5\) kpc, which sets a maximum density that can be resolved. The non-spherical shape of the halo is also clearly evident by inspecting the plots with different observer positions.

Recent estimates for the shape of the Milky Way’s halo (see, e.g., [37] and references therein), suggest that it may be flattened with a short/long axis ratio of 0.5. An independent estimate from the orbit of the Sagittarius debris stars yields a nearly spherical dark matter halo [28]. The simulated halo that we have chosen to analyse represents a typical prolate CDM halo with a short to long axis ratio of 0.4 and intermediate to long axis ratio of 0.5.

It is straightforward to estimate the effects of flattened dark matter halos by modifying Eq. (1) to accommodate triaxial shaped bodies. The simplest way to achieve this is to change from spherical coordinate \(r\) to

\[
\xi^2 = \frac{x^2}{a^2} + \frac{y^2}{b^2} + \frac{z^2}{c^2}
\]

where \(b > c\) for the oblate case, and \(b < c\) for prolate and we leave \(z\) as the axis of symmetry. A 2d visualization of these 3d shapes is illustrated in Fig. 4.

In Fig. 3 we plot spherical, oblate (2:1) and prolate (2:1) versions of the integral in Eq. (1) using the Moore et al. (1999) density profile. The observer is located on a plane parallel to the axis of symmetry, again at a distance \(R_o = 8.5\text{kpc}\) from the center of the halo. The halo shape leads to little difference towards the Galactic center, but at the anti-center prolate halos can be 100 times brighter than oblate halos.

We can also compare the predicted angular flux with that measured directly from the N-body simulation. The annihilation flux is averaged in ten degree bins from the simulated dark matter halo, along a great circle from the galactic center to its anti-center. This direct measurement of the flux is also plotted (as points) in Fig. 3. These data are particularly noisy due to the large numbers of substructure clumps in the simulation - the spike at \(ψ = 125°\) is due to a massive dark clump that happens to lie exactly along this chosen line of sight.

**III. SUBSTRUCTURE**
A. Enhancement of global flux due to substructure

Cold dark matter substructure clumps have singular density profiles that will be a significant source of annihilation products. The velocities and spatial distribution of dark matter substructure is unbiased with respect to the smooth dark matter background. Therefore, to first order, substructure increases the galactic sky brightness in any given direction. However, the details depend on how much substructure survives within the solar radius and also on how far down the mass function substructure halos form and survive.

First we will estimate the annihilation flux from clumps of dark matter that are known to exist in the Galactic halo i.e. the dark matter halos that surround the Magellanic Clouds and dwarf spheroidal galaxies. In fact, high-energy gamma-ray emission from the Large Magellanic Cloud (LMC) was detected with EGRET by Sreekumar et al. in 1992 (although the origin of this emission was reported to be the interaction of cosmic rays with interstellar matter).

We estimate the average flux, $\Phi_{AV}$, from the dark matter halos that surround some of the principal structures in the Local Group: The “Andromeda Galaxy”; M31 ($v_{peak} = 200$ km s$^{-1}$ at a distance of 700 kpc), The Large and Small Magellanic Clouds ($v_{peak} = 70$ km s$^{-1}$ and $v_{peak} = 40$ km s$^{-1}$, respectively; both at a distance of 50 kpc), Draco ($v_{peak} = 10$ km s$^{-1}$ at a distance of 50 kpc) and a small dark matter clump ($v_{peak} = 2$ km s$^{-1}$ at a distance of 10 kpc). A sketch of the geometry is given in Fig. 1.

The total flux from a substructure halo at distance $R_c$ from the observer is

$$\Phi_{TOT}(R_c) = \frac{K}{R_c^2} \int \rho^2(r) r^2 dr. \tag{6}$$

by considering the central $\Delta \Omega = 1^\circ \times 1^\circ$ patch over each clump, we define the maximum integration limit in above and the average flux is then

$$\Phi_{AV} = \frac{\Phi_{TOT}}{\Delta \Omega}. \tag{7}$$

(we set $\Delta \Omega$ in steradians, so we may compare directly with the smooth flux of Section 3).

For the dark matter distribution within the substructure clumps we use the Moore et al. profile, which provides a good fit to the smallest, well-resolved substructure halos. The concentration of CDM halos is a function of mass and for the density profile in Eq. 3 this can be written:

$$c_{moore} \approx 10^{2} \left( \frac{M_{vir}}{1h^{-1}M_{\odot}} \right)^{-0.084}. \tag{8}$$

This defines the scale radius of each substructure clump: $a_{M31} = 33.3$ kpc, $a_{LMC} = 6.7$ kpc, $a_{SMC} = 3.1$ kpc, $a_{Draco} = 0.5$ kpc, $a_{Tiny} = 0.05$ kpc.

The integral in Eq. 8 diverges as $r \to 0$ for the density profile that we are using, however, even the smallest substructure halos will have a maximum density set by the radius within which most of the neutralinos would have self annihilated. We therefore present results for the average flux from these clumps as a function of the minimum integration radius $R_{min}/a$ in Fig. 3 where $a$ is the scale radius as defined above.

For comparison, we plot the range of background emission at the Galactic anti-center as the shaded line in Fig. 3. The tiny clump is only marginally visible above the background flux (depending on whether or not the Galactic halo is prolate or oblate) whereas most of the subhalos are easily visible. Also for comparison we have plotted the flux from the inner region of the galaxy which is the brightest of these sources.

Although the Galactic halo is expected to contain just a few clumps more massive that the Magellanic Clouds, there are many thousands of smaller mass objects. The mass function of substructure is a power law close to $dn(m)/dm \propto m^{-1.9}$ or in terms of circular velocity $dn(v_c)/dv_c \propto v_c^{-3.8}$ (34). Above a circular velocity $v_{peak} = 10$ km s$^{-1}$ and 1 km s$^{-1}$ we expect the galactic halo to host roughly 1000 and $5 \times 10^5$ substructure halos respectively. Future simulations should be able to measure how far down the mass function substructure halos can survive as well as to determine their central density profiles. (We note that the highest resolution simulation to date resolved the substructure within a dark matter mini-halo of mass $10^7 M_{\odot}$. The force resolution was 10 parsecs and the mass resolution was 10M$_{\odot}$ allowing substructure with peak circular velocities as low as a few hundred meters per second to be resolved. The survival of substructure continues even down to this scale, where the slope of the power spectrum is close to -3.)

We calculate the total flux from substructure using Monte-Carlo techniques. First we generate a list of peak circular velocities and positions of $5 \times 10^5$ substructure halos in the range of 1–70 km s$^{-1}$. (Distances are randomly selected using the Moore et al. density profile and peak circular velocities are randomly assigned from a power law distribution scaling as $v^{-3.8}$). For each lump, we estimate its total flux as in the previous cases, integrating Eq. 8 with a density profile scaled according to Eq. 8 for the concentration.

In the absence of further constraints on the possible value for $R_{min}$, we use the same criteria as before and choose it to be a fixed fraction of the virial radius, $R_{min} \approx 10^{-12} r_{200}$. This corresponds to a mean density of $\approx 10^{22} M_{\odot}$ kpc$^{-3}$ for the galactic halo. The results are not too sensitive to the value of the minimum integration radius as is apparent from inspection of Fig. 3. The total flux is then averaged over the entire sky and we repeat this process in order to estimate the variance. The cumulative distribution of flux above a given substructure peak circular velocity ($\Sigma \Phi_{TOT}$) is plotted for ten of these random halo realisations in Fig. 3.

It is evident from this plot that the effects of includ-
ing the entire mass spectrum of substructure is quite
dramatic and boosts the expected flux from the smooth
halo by several orders of magnitude. However, most of
the flux arises from the subhalos with circular velocities
larger than 10 km s$^{-1}$. Extrapolating to very small halos
would not change the total flux by a large factor.

To quantify the brightening of the background due to
substructure, we have calculate the average flux due to
all clumps with $v_{\text{peak}} > 1$ km s$^{-1}$ within a spherical halo.
The point plotted in Fig. 3 represents this contribution
to the flux, where the error-bar is the $1\sigma$ variation among
the different Monte-Carlo models. From this plot we see
that the flux due to substructure is over two orders of
magnitude larger than the smooth background from
a spherical halo. We note that one needs to observe a
fairly large fraction of the sky ($\approx 1$) to ensure that the flux due to substructure is over two orders of
magnitude brighter than the smooth background from
a spherical halo. We note that one needs to observe a
fairly large fraction of the sky ($\approx 100$ square degrees) to
ensure a significant number of clumps lie in the field of
view. (Also note that the variance at high peak circular
velocities is due to the proximity of the largest few
dark matter substructures, however, the mean total flux
converges to similar values for each Monte-Carlo model.)

B. The flux due to substructure in prolate and
oblate halos

Not only is the mean flux at a given position on the
sky dominated by substructure halos, the spatial distri-
bution of flux across the sky will be determined by the
substructure. The convergence study by Ghigna et al.
(Ref. [34]) shows that substructure halos trace the global
mass distribution of the halo. Therefore, we can use the
N-body simulations to generate Monte-Carlo distribu-
tions of substructure halos and construct all-sky maps of
the expected gamma-ray flux. We take a random particle
from the simulation and assign a circular velocity from
a distribution $dn(v_c)/dv_c \propto v_c^{-3.8}$. For each sub-halo we
calculate its total annihilation flux and then repeat the
process until we have 500,000 halos above a circular ve-
locity of 1 km s$^{-1}$.

Figure 3(a) and Fig. 3(b) show the resulting sky distri-
bution of flux from sub-halos binned in one degree bins
where the observer has been placed in the short and long
axis of the simulation respectively. Large substructure
halos, such as the Magellanic Clouds in our own halo, will
contain its own gravitationally bound sub-halos which
leads to clustering of gamma-ray emission in the all sky
maps.

Future observations may only be able to make deep
strip maps therefore in Fig. 11 we have binned the flux
along lines of constant galactic $l$ and $b$ with the observer
placed in the short and long axis of the global density dis-
tribution. From these plots we can see that the emission
from substructure peaks at the galactic center, as one
would expect, this effect is not that different for spheri-
cal halos than for prolate or oblate halos.

Since the substructure traces the global mass distribu-
tion, a prolate halo would also have a prolate distribu-
tion of satellites. Therefore we can study the variation of flux
within smooth prolate or oblate halos to examine how the
background flux from substructure can be used to quan-
tify the halo shape. We calculate the observed flux as a
function $l$ and $b$ for spherical, prolate and oblate flattened
2:1 and 3:1 geometries. In each case, the density profile
is taken from Eq. 3 and again the observer is placed in
the short (Figure 11(a)) or long (Figure 11(b)) axis of
symmetry.

These plots show how the distribution of flux on the
sky can vary significantly depending on the shape of the
density distribution and on where the observer is situated
within the halo.

C. The distribution of point sources

Individual substructures may be observed and quanti-
fied if the resolution of the telescope is sufficient. How-
ever, all of the past and present observations would only
detect substructure as unresolved point sources. The dis-
tribution of their fluxes (and spatial distribution on the
sky) may be used to rule out alternative origins, such as
extra-galactic sources. In Fig. 12 we plot the cumulative
distribution of point sources above a given flux within
one degree square bins. The two curves consider sub-
structure with peak circular velocities larger than 10 km
s$^{-1}$ and 1 km s$^{-1}$. The number density of the brightest
sources in the sky scales as $N \propto F^{-0.7}$.

Higher resolution simulations are vital to quantify how
much substructure survives within the galactic halo, how
it is spatially distributed and to quantify the internal
structure of surviving substructure. However, Fig. 12
gives an idea of what to expect if an all sky survey is
carried out that is capable of detecting the brightest sub-
structure halos.

IV. CONCLUSIONS

Numerical simulations that follow the growth of struc-
ture within a universe dominated by neutralinos (cold
dark matter) have achieved a resolution that allows their
global structure and internal structure to be quantified.
The density profiles, shapes of dark matter halos, abun-
dance and properties of dark matter substructure, all
play an important role in determining the absolute sur-
fase brightness of observable products from dark matter
annihilation.

We have used the results from the highest resolu-
tion simulations ever performed of CDM halos to ex-
amine the expected all-sky distribution of gamma-rays
from neutralino annihilation. Substructure can boost
the expected flux significantly over that originating from
a smooth dark matter halo. Thus, gamma-ray obser-
vations, such as EGRET data, may already have the
potential of constraining a large part of the parameter
range of the neutralino cross-sections. The distinguishing shapes of CDM halos and the unique spatial and flux distribution of point sources from substructure within the Galactic halo should allow a unique identification of observational data with dark matter.

ACKNOWLEDGMENTS

The authors would like to thank Prof. Arnold Wolfendale for numerous discussions and suggestions that have improved the quality of this work. Carlos Calcáneo–Roldán continues his research thanks to the generous support from the People of México, through a grant by CONACyT. Ben Moore is a Royal Society research fellow. Computations were carried out as part of the Virgo consortium.

[1] J. Ellis, invited talk presented at the nobel symposium. Haga Slott, Sweden. August 1998. [astro-ph/9812211].
[2] M. Davis, G. Efstathiou, C. S. Frenk & S. D. M. White, Astrophys. J., 292, 371. (1985).
[3] C. M. Baugh, S Cole, C. S. Frenk, C. G. Lacey, Astrophys. J., 498, 521 (1998).
[4] B. Moore, Nature 370 629 (1994).
[5] B. Moore, S. Ghigna, F. Governato, G. Lake, T. Quinn, J. Stadel & P. Tozzi, Astrophys. J. Lett. 524, L19 (1999).
[6] A. A. Klypin, A. V. Kravtsov, O. Valenzuela & F. Prada, Astrophys. J. 522, 82 (1999).
[7] J. Collar, Phys. Rev. D 59, 063514 (1999).
[8] J. E. Gunn, B. W. Lee, I. Lerche, D. N. Schramm, & G. Steigman, Astrophys. J. 223, 1015 (1978).
[9] F. W. Stecker, Astrophys. J. 223, 1032 (1978).
[10] J. Silk, & M. Srednicki, Phys. Rev. Lett. 53, 624 (1984).
[11] M. S. Turner, Phys. Rev. D 34, 061921 (1986).
[12] J. Silk, & H. Bloemen, Astrophys. J. 313, L47 (1987).
[13] A. Bouquet, P. Salati, J. Silk, Phys. Rev. D 40, 103168 (1989).
[14] G. Lake, Nature 346, 39 (1990).
[15] L. Bergström, J. Edsjö, P. Gondolo, Phys. Rev. D 58, 103519 (1998).
[16] L. Bergström, P. Ullio, J. H. Buckley, Astropart. Phys. 9, 137 (1998).
[17] E. A. Baltz, & J. Edsjö, Phys. Rev. D 59, 023511 (1999).
[18] P. Gondolo, & J. Silk, Phys. Rev. Lett. 83, 1719 (1999).
[19] E. A. Baltz, C. Briot, P. Salati, R. Taillet, & J. Silk, Phys. Rev. D 61, 023514 (2000).
[20] D. D. Dixon, D. H. Hartman, E. D. Kolaczyk, J. Samimi, New Astron. 3, 539 (1998).
[21] H. A. Mayer-Hasselwander, D. L. Bertsch, B. L. Dingus, A. Eckart, J. A. Esposito, R. Genzel, R. C. Hartman, S. D. Hunter, G. Kanbach, D. A. Kniffen, Y. C. Lin, P. F. Michelson, A. Mücke, C. von Montigny, R. Mukherjee, P. L. Nolan, M. Pohl, O. Reimer, D. A. Sreekumar, D. J. Thompson, Astron. Astrophys. 335 161 (1998).
[22] R.C. Hartman, D. L. Bertsch, S. D. Bloom, A. W. Chen, P. Deines-Jones, J. A. Esposito, C. E. Fichtel, D. P. Friedlander, S. D. Hunter, L. M. McDonald, P. Sreekumar, D. J. Thompson, B. B. Jones, Y. C. Lin, P. F. Michelson, P. L. Nolan, W. F. Tompkins, G. Kanbach, H. A. Mayer-Hasselwander, A. Mücke, M. Pohl, O. Reimer, D. A. Kniffen, E. J. Schneid, C. von Montigny, R. Mukherjee, B. L. Dingus, Astrophys. J. Supp. 123, 79 (1999).
[23] N. Gehrels, P. Michelson, Astropart. Phys. 11, 277 (1999).
[24] R. W. Lessard, Astropart. Phys. 11, 243 (1999).
[25] R. G. Carlberg, Astrophys. J. 433, 468 (1994).
[26] J. F. Navarro, C. S. Frenk, & S. D. M. White, Astrophys. J. 462, 563 (1996).
[27] B. Moore, F. Governato, N. J. Stadel, T. Quinn & G. Lake, Astrophys. J. Lett. 499, L5 (1998).
[28] S. Ghigna, B. Moore, F. Governato, G. Lake, T. Quinn, & J. Stadel, Mon. Not. R. Astron. Soc. 300, 446 (1998).
[29] L. Bergström, J. Edsjö, P. Gondolo, P. Ullio Phys. Rev. D 59, 043506 (1999).
[30] G. Jungman, M. Kamionkowski, & K. Griest, Phys. Rep. 267, 195 (1996).
[31] L. Bergström, Nucl. Phys. B – Proc. Suppl. 70, 31 (1999).
[32] F. J. Kerr, & D. Lynden-Bell, Mon. Not. R. Astron. Soc. 221, 1023 (1986).
[33] Y.P. Jing & Y. Suto, Astrophys. J. 529, L69 (2000).
[34] J. Silk, B. Moore, F. Governato., G. Lake, T. Quinn & J. Stadel, Astrophys. J. in press (2000).
[35] A. V. Kravtsov, A. A. Klypin, J. S. Bullock, & J. R. Primack, Astrophys. J. 502, 48 (1998).
[36] J. Barnes, & G. Efstathiou, Astrophys. J. 319, 575 (1987).
[37] R. P. Olling, & M. R. Merrifield, ASP Conference Series 182 (1999).
[38] R. Ibata, G.F. Lewis, M. Irwin, E. Totten & T. Quinn, Astrophys. J. submitted (2000), [astro-ph/0004011].
[39] P. Sreekumar, D. L. Bertsch, B. L. Dingus, C. E. Fichtel, R. C. Hartman, S. D. Hunter, G. Kanbach, D. A. Kniffen, Y. C. Lin, J. R. Mattox, H. A. Mayer-Hasselwander, P. F. Michelson, C. von Montigny, P. L. Nolan, K. Pinkau, E. J. Schneid & D. J. Thompson, Astrophys. J. Lett. 400, L67 (1992).
[40] A. Klypin, S. Gottlöber & A. V. Kravtsov, Astrophys. J. 516, 530 (1999).

**FIG. 1.** The circular velocity curves $V_c(r) = \sqrt{GM(r)/r}$, and (b) density profiles are plotted as a function of the radius for each of the halo models considered in the text.

**FIG. 2.** The gamma ray flux from neutralino annihilation, $\phi(\psi)$, plotted as a function of the angular distance from the galactic center $\psi$. The curves show the results using the three different density profiles plotted in Fig. 1. The flux at a given position is averaged over $4\pi$ steradians.
FIG. 3. All-sky maps of the gamma ray background constructed using a single high-resolution N-body simulation of a cold dark matter halo. The observer has been placed in the short (a) and long (b) axis of the simulated halo.

FIG. 4. The left panel shows a unit oblate ellipsoid and the right hand panel shows a unit prolate ellipsoid. The axial ratios for both are 2:1.

FIG. 5. The gamma ray flux, $\phi$, plotted as a function of angle $\psi$, for smooth halos of the same total mass using the density profile given in Eq. 2 for spherical, oblate and prolate halo geometries. The points are values of the flux measured directly from the N-body halo illustrated in Fig. 3.

FIG. 6. A sketch showing the geometry of an observer in the galaxy viewing substructure in the galactic halo.

FIG. 7. The gamma ray flux, $\Phi_{AV}$, plotted as a function of minimum integration radius $R_{min}$ for halo substructure of different circular velocities and distances as detailed in the text. The shaded region shows the range of background values at the Galactic anti-center that can be expected depending on the halo shape. The point is the average flux due to all clumps with $v_{\text{peak}} \geq 1$ km s$^{-1}$. Note that the size of the error bar on this point depends on the area of the sky surveyed.

FIG. 8. The cumulative gamma-ray flux from halo substructures, $\Sigma \phi_{TOT}(v > v_{\text{peak}})$, above a given substructure circular velocity $v_{\text{peak}}$. The ten different curves correspond to different Monte-Carlo realizations of a Galactic halo of substructure halos. The flux is averaged over $4\pi$ steradian and can be compared with the flux from the smooth halo from Fig. 2 and Fig. 5.

FIG. 9. All-sky map of the gamma ray background that arises solely from dark matter substructures. The positions and circular velocities of sub-halos above a circular velocity of 1 kms$^{-1}$ are drawn from the N-body simulations but the flux from each halo is calculated analytically. The observer is located on the short (a) and long (b) axis of symmetry. The grey scale corresponds to the log of the flux of annihilation products.

FIG. 10. The average gamma-ray flux per square degree from dark matter substructure as measured within the simulated CDM halo along a great circle of constant galactic latitude (a) and longitude (b). The average has been taken over a strip of width 44 degrees, the left hand plot represents the view along the short axis while the right hand side is the view along the long axis.

FIG. 11. The effect of halo shape on the gamma-ray flux. Halo density profiles are drawn from spherical, oblate or prolate distributions with the indicated axis ratios. The observer is placed in the short axis (a) whilst the in (b) the observer is in the long axis.

FIG. 12. The cumulative number of gamma-ray sources above a given flux within a window $\Delta \Omega = 1^\circ \times 1^\circ$. The two curves are for substructure halos with circular velocities larger than 10 kms$^{-1}$ (dashed line) and 1 kms$^{-1}$ (solid line).
