REVIEW

Models of very-low-mass stars, brown dwarfs and exoplanets

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Within the next few years, GAIA and several instruments aiming to image extrasolar planets will be ready. In parallel, low-mass planets are being sought around red dwarfs, which offer more favourable conditions, for both radial velocity detection and transit studies, than solar-type stars. In this paper, the authors of a model atmosphere code that has allowed the detection of water vapour in the atmosphere of hot Jupiters review recent advances in modelling the stellar to substellar transition. The revised solar oxygen abundances and cloud model allow the photometric and spectroscopic properties of this transition to be reproduced for the first time. Also presented are highlight results of a model atmosphere grid for stars, brown dwarfs and extrasolar planets.

Keywords: PHOENIX; CO5BOLD; very-low-mass stars; brown dwarfs; stars; opacities

1. Introduction

Since the spectroscopic observations of very-low-mass stars (VLMs, late 1980s), brown dwarfs (mid-1990s) and extrasolar planets (mid-2000s) have become available, the MK spectral classification has had to be extended beyond K and M to the newly defined classes L and T. One of the most important challenges in modelling their atmospheres and spectroscopic properties has been high-temperature molecular opacities and cloud formation. The K dwarfs show the onset of formation of metal hydrides (starting around $T_{\text{eff}} \sim 4500\,\text{K}$), TiO and CO (below $T_{\text{eff}} \sim 4000\,\text{K}$), while water vapour forms in early M dwarfs ($T_{\text{eff}} \sim 3900–2000\,\text{K}$), and methane, ammonia and carbon dioxide are detected in late-type brown dwarfs ($T_{\text{eff}} \sim 300–1600\,\text{K}$) and in extrasolar giant planets. The latter are observed either by transit ($T_{\text{eff}} \sim 1000–2000\,\text{K}$, depending on the spectral type of the central star and the distance to the star) or by imaging (young planets of $T_{\text{eff}} \sim 300–2000\,\text{K}$, depending on their mass and age).

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One contribution of 17 to a Theo Murphy Meeting Issue ‘Water in the gas phase’.
The modelling of the atmospheres of VLMs has evolved (as illustrated here with the development of the PHOENIX atmosphere code, which has allowed the detection of water vapour in the atmospheres of extrasolar planets by Barman et al. [1,2]) with the extension of computing capacities from an analytical treatment of the transfer equation using moments of the radiation field [3], to a line-by-line opacity sampling in spherical symmetry [4–6] and, more recently, to three-dimensional radiation transfer [7]. In parallel with detailed radiative transfer in an assumed static environment, hydrodynamical simulations have been developed to reach a realistic representation of the granulation and its induced line shifts for the Sun and Sun-like stars [8] by using a non-grey (multi-group binning of opacities) radiative transfer with a pure blackbody source function (scattering is neglected).

2. Molecular opacities

While earlier work was developed for the study of red giant stars, the pioneering work on the modelling of VLM atmospheres has been provided [3,9,10] using
a band model or just overlapping line approximation opacities developed by Kivel et al. [11] and adapted for astrophysical use by Golden [12]. More realistic model atmospheres and synthetic spectra for VLMs, brown dwarfs and extrasolar planets have been made possible thanks to the development of accurate opacities calculated, often \textit{ab initio}, for atmospheric layers where temperatures can reach 3000 K. The process of improvement has been especially remarkable in the case of water vapour line lists. Indeed, water vapour has seen an important evolution through the years, from band model approximations to straight means based on hot flame experiments, and then to \textit{ab initio} computations. Nevertheless, the atmosphere models have failed to reproduce the strength of the water bands that shape the low-resolution ($R \leq 300$) infrared (IR) spectral energy distributions (SEDs) of M dwarfs. At the lower temperatures of brown dwarfs, methane and ammonia rival the effect of water. Therefore, the discrepancies in the model synthetic spectra were believed to be due to inaccurate or incomplete molecular opacities. In particular, water vapour was suspected because the discrepancies were observed at IR wavelengths in the relative brightnesses of the flux peaks between water vapour bands. As can be seen from figure 1, where the models are compared with the IR spectrum of the M8e dwarf VB10, the water vapour opacity profile that shapes this part of the spectrum has changed strongly over time with the improvement of computational capacities and better knowledge of the interaction potential surface. The most recent \textit{ab initio} results confirm the earliest hot flame laboratory experimental results by Ludwig [14]. However, in general, most opacity profiles produce an excess opacity (or lack of flux in the model) in the $K$ bandpass. Only the UCL1994 line list (owing to incompleteness, and with many of its deviations cancelling out over the bandpasses) could produce seemingly correct $J - K_s$ colours.

3. The revised solar abundances

Model atmospheres for VLMs and, in general, for other stars assume scaled solar abundances for all heavy elements, with some enrichment of $\alpha$-process elements (the result of ‘pollution’ of the star-forming gas by the explosion of a supernova) when appropriate in the case of metal-depleted subdwarfs of the Galactic thick disc, halo and globular clusters. The revision of the solar abundances based on radiation hydrodynamical (RHD) simulations of the solar atmosphere, on improvements in the quality of the spectroscopic observations of the Sun, and on its detailed line profile analysis by two separate groups using independent hydro codes and spectral synthesis codes [21,22] yield an oxygen reduction of 0.11–0.19 dex (up to 34\%) compared with the previously used abundances of Grevesse et al. [23]. Since the overall SED of late K dwarfs, M dwarfs, brown dwarfs and exoplanets is governed by oxygen compounds (TiO and VO in the optical, and water vapour and CO in the IR), the input elemental oxygen abundance used in the equation of state is of major importance. Figure 2 shows an example of these effects for the optical and IR SED of the M5.5 dwarf system Gl866. However, at other effective temperatures, even stronger photometric effects can be seen, where the near-IR SEDs of different models diverge more (figure 3). The comparison shows significant improvement when compared with the older models shown in figure 1, except for excess flux in the $H$ bandpass near 1.7 $\mu$m due to incomplete...
Figure 2. A BT-Settl synthetic spectrum with log $g = 5.0$ and solar metallicity by Asplund et al. [21] (thin grey/green full line; [M/H] = 0.0) compared with the combined SED of the red dwarf triple system Gl 866 [24,25]. The observations of GJ 866ABC were combined from a Mt. Stromlo optical spectrum (M. Bessell 2009, private communication) and SpeX IR spectrum taken at the NASA IRTF [26] (thick black curve). For comparison a model using the same parameters and physical setup with the Grevesse et al. [23] abundances is also shown (grey/blue dashed line). The models have been scaled to the observed absolute flux assuming two equal $T_{\text{eff}} = 2920$ K components of 0.157 $R_{\odot}$ (solar radii) and a third with $T_{\text{eff}} = 2700$ K and 0.126 $R_{\odot}$. (Online version in colour.)

FeH opacity data for this region. The comparison has particularly improved in the Wing Ford band of FeH near 0.99 μm, and in the VO bands thanks to line lists provided by B. Plez (GRAAL, Montpellier, France), although inaccurate or incomplete opacities still affect the models at optical wavelengths (e.g. the TiO line list by Langhoff [32]).

Figure 3 compares the theoretical isochrones (assuming an age of 5 Gyr) with the $T_{\text{eff}}$ estimates [27] and reveals that the NextGen models [4–6] systematically and increasingly overestimate $T_{\text{eff}}$ through the lower main sequence, while the
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Grevesse et al. [23]:
- NextGen
- AMES-Cond
- AMES-Dusty
- BT-NextGen
- BT-Dusty
Asplund et al. [21]:
- BT-Dusty
- BT-Settl

Figure 3. The estimated $T_{\text{eff}}$ for M dwarfs by Casagrande et al. [27] and for brown dwarfs by Golimowski et al. [28] and Vrba et al. [29] are compared with the NextGen isochrones for 5 Gyr [30,31] using various generations of model atmospheres: NextGen (thick black line), the limiting AMES-Cond/Dusty cases by Allard et al. [19] (dotted blue and dashed red lines), and the current BT-Settl models using the Asplund et al. [21] solar abundances (full green line). The Gl 866 system fitted in figure 2 is highlighted by darker colours and shown with its relatively large photometric error bars at $J - K_s = 0.9$. (Online version in colour.)

AMES-Cond/Dusty [19] models, on the contrary, underestimate $T_{\text{eff}}$ as a function of $J - K_s$ colour. This situation is relieved when using the current models (labelled BT-Settl in the figure) based on the revised solar abundances, and the models now agree fairly well with most of the empirical estimations of $T_{\text{eff}}$. The current model atmospheres have not yet been used as surface boundary conditions to interior and evolution calculations, and simply provide the synthetic colour tables interpolated on the published theoretical isochrones [31]. Even if the atmospheres partly control the cooling and evolution of M dwarfs [33], differences introduced in the surface boundary conditions by changes in the model atmosphere composition have a negligible effect.

4. Cloud formation

One of the most important challenges in modelling these atmospheres (below 2600 K) is the formation of clouds. Tsuji et al. [34] identified dust formation by recognizing the condensation temperatures of hot dust grains (enstatite MgSiO$_3$,...
forsterite $\text{Mg}_2\text{SiO}_4$ and corundum $\text{Al}_2\text{O}_3$ crystals) that occur in the line-forming layers ($\tau \approx 10^{-4}$ to $10^{-2}$) of their atmospheres. The cloud composition, according to equilibrium chemistry, goes from zirconium oxide ($\text{ZrO}_2$), to refractory ceramics (perovskite $\text{CaTiO}_3$ and corundum $\text{Al}_2\text{O}_3$), to silicates (e.g., forsterite $\text{Mg}_2\text{SiO}_4$), to salts ($\text{CsCl, RbCl, NaCl}$) and finally to ices ($\text{H}_2\text{O, NH}_3, \text{NH}_4\text{SH}$) as brown dwarfs cool down over time from M through L, T and Y spectral types [19,35]. This assumed (by Allard et al. [19]) that sub-micrometre-sized crystal formation causes the weakening and vanishing of TiO and VO molecular bands (via $\text{CaTiO}_3$, $\text{TiO}_2$ and $\text{VO}_2$ grains) from the optical spectra of late M and L dwarfs, revealing CrH and FeH bands otherwise hidden by the molecular pseudo-continuum, and the resonance doublets of alkali transitions, which only condense onto salts in late T dwarfs. The scattering effect of this fine dust is Rayleigh scattering, which provides veiling to the optical SED of late M and L dwarfs, while the greenhouse effect due to the dust cloud causes their IR colours to become extremely red when compared with those of hotter low-mass stars. The upper atmosphere, above the cloud layers, is depleted of condensible material and significantly cooled down by the reduced or missing pseudo-continuum opacities.

One common approach has been to explore the limiting properties of cloud formation. One limit is the case where sedimentation or gravitational settling is assumed to be fully efficient, such as case B of Tsuji [36], the AMES-Cond or condensed phase models of Allard et al. [19], the clear case of Ackerman & Marley [37] and the cloud-free case of Burrows et al. [38]. The other limit is the case where gravitational settling is assumed to be inefficient and dust, often only forsterite, forms in equilibrium with the gas phase, such as case A of Tsuji [36], the AMES-Dusty or dusty models of Allard et al. [19], the cloudy case of Ackerman & Marley [37] or case B of Burrows et al. [38]. These limiting cases of maximum dust content agree in describing the evolution of brown dwarfs from a molecular opacity-governed SED towards a blackbody SED below 1500 K. This description was suitable, at least in the case of the AMES-Dusty models, in reproducing the IR colours of L dwarfs. The cloud-free limiting case, on the other hand, allowed the colours of T dwarfs to be reproduced to some degree. Figure 4 shows this situation for the AMES-Cond/Dusty limiting case models of Allard et al. [19] compared with the effective temperature estimates obtained by integration of the observed SED [28,29].

The purpose of a cloud model is therefore to go beyond these limiting cases and define the number density and size distribution of condensates as functions of depth in the atmosphere. The discovery of dust clouds in M dwarfs and brown dwarfs has therefore triggered the development of cloud models building upon the pioneering work in the context of planetary atmospheres developed in the earlier studies [39–41]. The Lewis model is an updraft model (considering that condensation occurs in a gas bubble that is advected from deeper layers). Owing to lack of knowledge of the velocity field and diffusion coefficient of condensates in the atmospheres of the planets of the Solar System, Lewis [39] simply assumed that the advection velocity is equal to the sedimentation velocity, thereby preserving condensible material in the condensation layers. This cloud model did not account for grain sizes. Rossow [40], on the other hand, developed characteristic time scales as a function of particle size for the main microphysical processes involved (condensation, coagulation, coalescence and sedimentation). The intersections of these characteristic time scales give an
Figure 4. Same plot as figure 3 but zooming out and extending into the brown dwarf region of the diagram. This region below 2500 K is dominated by dust formation (essentially forsterite and other silicates). The AMES-Cond/Dusty model atmosphere limiting cases provide a description of the span in colours of the brown dwarfs in this diagram for a given age (here 5 Gyr). The BT-Settl models succeed in explaining even the most extreme colours of brown dwarfs. (Online version in colour.)

estimate of the condensate number densities and mean grain sizes. However, this model made several explicit assumptions concerning the efficiency of supersaturation, coagulation, etc.

Helling et al. [42] have compared different cloud models and their impact on model atmospheres. Most cloud models define the cloud base as the evaporation layer provided by the equilibrium chemistry. In the unified cloud models of Tsuji and co-workers [36, 43], a parametrization of the radial location of the cloud top by way of an adjustable parameter $T_{\text{crit}}$ was used. This choice permits the cloud extension effects on the spectra of these objects to be determined but does not allow the stellar–substellar transition to be reproduced with a unique value of $T_{\text{crit}}$, as the cloud extension depends on the atmospheric parameters.

Allard et al. [44], using PHOENIX and the index of refraction of up to 40 condensible species, have applied the Rossow cloud model, ignoring coalescence and coagulation, and comparing the time scales of condensation, sedimentation and mixing (extrapolated from the convective velocities into the convectively stable layers), and assuming efficient nucleation (monomer equilibrium densities). The cloud model was then solved layer by layer inside out to account for the sequence of formation of grain species as a function of cooling of the gas. However,
this version of the BT-Settl (with gravitational settling) models did not allow the formation of enough dust in brown dwarf atmospheres, owing to a very conservative prescribed supersaturation value.

Ackerman & Marley [37] have solved the particle diffusion problem of condensates by assuming a parametrized sedimentation efficiency $f_{\text{sed}}$ (constant through the atmosphere) and a mixing assumed constant and fixed to its maximum value (maximum of the inner convection zone). Saumon & Marley [45] found that their models could not produce the colour change with a single value of $f_{\text{sed}}$.

Helling et al. [46] used the PHOENIX code to compute the Drift–Phoenix models. The cloud model used, in contrast to all other cases mentioned, studies the nucleation and growth of grains as they sediment down into the atmosphere. This cloud model determines the number density and size distribution of grains by one-dimensional nucleation simulations, and the resulting distribution is read in by PHOENIX, which computes the resulting opacities and radiative transfer. These models solve the nucleation problem, but only for the assumed monomer types, and have been successfully applied to fit the dusty atmospheres of L dwarfs, but the reversal in IR colours observed for the L–T transition could not be explained [47].

However, none of these models treated the mixing properties of the atmosphere and the resulting diffusion mechanism realistically enough to reproduce the brown dwarf spectral transition without changing the cloud parameters. Freytag et al. [48] have therefore addressed the complementary though important issue of mixing and diffusion in these atmospheres by two-dimensional RHD simulations, using the PHOENIX gas opacities in a multi-group opacity scheme and forsterite with geometric cross sections. These simulations assume efficient nucleation, using monomer densities estimated from the total available density of silicon (least abundant element in the solar composition involved in forsterite). They found that gravity waves play a decisive role in cloud formation, while around $T_{\text{eff}} \leq 2200$ K the cloud layers become optically thick enough to initiate cloud convection, which participates in the mixing. Overshoot can also be important in the deepest layers.

These RHD simulations allow an estimation of the diffusion processes that bring fresh condensible material from the hotter lower layers to the cloud-forming layers. We have therefore updated our cloud model (BT-Settl models) to account for the mixing prescribed by the RHD simulations. Another important improvement concerns the supersaturation, which has been computed rather than using the fixed conservative value recommended by Rossow. One can see from figure 4 that the late-type M and early-type L dwarfs behave as if dust is formed nearly in equilibrium with the gas phase, with extremely red colours in some agreement with the BT-Dusty models. The BT-Settl models reproduce the main sequence down to the L-type brown dwarf regime, subjected in the $K$ bandpass to the greenhouse effect of dust clouds, before turning to the blue in the late L and T dwarf regime as a result of methane formation in the $K$ bandpass. This constitutes a major improvement over previous models, and shows promise that, in the near future, we can reach a clear explanation of the stellar–substellar transition.

Diffusion has also been held responsible for deviations in ultracool atmospheres from gas-phase chemical equilibrium, as noted in early observations of T dwarfs showing an excess of carbon monoxide absorption [49,50]. More recently, carbon
dioxide [51], which was not expected at such low temperatures, has been detected. Similarly, ammonia has been shown to be underabundant [52]. This is understood as the result of slowing down of crucial chemical reaction steps, so that some important molecules (CH$_4$, NH$_3$) would not have the time to form in equilibrium while undergoing mixing, whereas others (CO, CO$_2$, N$_2$) remain at enhanced abundances. The RHD simulations of Freytag et al. [48] have allowed the underlying mixing processes to be understood, obviating the need to describe them with an additional free parameter.

5. Applications to exoplanet science

Several IR integral field spectrographs combined with coronograph and adaptive optic instruments that are being developed will come online before 2013 (SPHERE at the Very Large Telescope (VLT), the Gemini Planet Imager at Gemini South, Project1640 at Mount Palomar, etc.). The 39m European Extremely Large Telescope (E-ELT) at Cerro Armazones in Chile due around 2020 will also be very ideally suited for planet imaging. The models developed for VLMs and brown dwarfs are a unique opportunity, if they can explain the stellar–substellar transition, to provide great support for the characterization of imaged exoplanets. We have therefore developed the BT-Settl model atmosphere grid to encompass the parameter regime of these objects (surface gravity around log $g = 4.0$, $T_{\text{eff}} < 2000$K).

These planets are typically found at several dozens of astronomical units (AU) from the star, and, since the observations are done in the IR, the non-irradiated models can even be used directly. Indeed, Barman et al. [53] have shown that the effects of radiation from a star impinging on the planetary atmosphere are Rayleigh scattering of the stellar light by H$_2$ molecules (or clouds, if present) at optical wavelengths (below 1$\mu$m for solar-type stars), while the impact on the interior and evolution properties becomes negligible for orbital distances exceeding 0.1 AU. Nevertheless, for 2012 we are developing irradiated models and the capacity to compute them via the PHOENIX simulator (see §6).

6. Summary and future prospects

We report progress on the development of a new model atmosphere grid for stars, brown dwarfs and young planets, named BT-Settl. It has been computed using the PHOENIX code updated for: (i) the line lists of water by Barber & Tennyson (BT2) [54], methane using the Spherical Top Data System (STDS) [55], ammonia [56] and CO$_2$ opacity from the Carbon Dioxide Spectroscopic Databank (CDSD-1000) [57], (ii) the solar abundances revised by Asplund et al. [21] and (iii) a cloud model accounting for more detailed supersaturation and RHD mixing. The grid covers the whole range of stars to young planets, 400 K $< T_{\text{eff}} < 70 000$ K, $-0.5 < \log g < 5.5$ and $-4.0 < [\text{M/H}] < +0.5$, including values of the $\alpha$-element enhancement (supernovae enrichment of the star-forming material) between $+0.0$ and $+0.6$. Models are available at the PHOENIX simulator website (http://phoenix.ens-lyon.fr/simulator/) and are in preparation for publication to

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serve, among others, the GAIA, MUSE and SPHERE/GPI/P1640 instruments due to come online in the near future. Corresponding evolution models are expected for 2012.

We found that the previously used NextGen models systematically overestimate $T_{\text{eff}}$ below 3500 K by as much as 500 K. The water vapour opacity profile has converged with the most recent line lists reproducing laboratory results, but could not explain this discrepancy. The solution came instead from the revision of the solar abundances, which changes the strength of the water vapour absorption bands, and therefore allows the reproduction of the spectroscopic and photometric properties of M dwarfs as late as M6. Later-type M dwarfs are affected by dust formation, and cloud modelling is important to understand their properties. We find that the Rossow cloud model allows, with revisions to the supersaturation and mixing, the stellar–substellar transition to be reproduced. A small offset persists, however, in the M–L transition. It is possible that all the current cloud models are not efficient enough in producing dust at the onset of the cloud formation regime. Detailed nucleation studies could allow this issue to be resolved in the future. Other uncertainties affect the current cloud modelling, such as the assumption of spherical non-porous grains, whereas grains form as fractals in the laboratory. Constraining the models therefore remains very important.

Beyond cloud modelling and molecular opacities, model atmospheres for these objects require reaction rates for the most abundant molecules and/or most important absorbers. Furthermore, these atmospheres are composed of molecular hydrogen, which constitutes the main source of collisions. Also needed therefore are collision rates (by $H_2$) and corresponding damping constants for the broadening molecular lines.

In order to say something about the spectral variability of VLMs, brown dwarfs and planets, three-dimensional global or ‘star-in-a-box’ RHD simulations with rotation will be required. This is our current project supported by the French ‘Agence Nationale de la Recherche’ for the period 2010–2015.

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References

1 Barman, T. 2007 Identification of absorption features in an extrasolar planet atmosphere. *Astrophys. J. Lett.* **661**, L191. (doi:10.1086/518736)

2 Barman, T. S. 2008 On the presence of water and global circulation in the transiting planet HD 189733b. *Astrophys. J. Lett.* **676**, L61. (doi:10.1086/587056)

3 Allard, F. 1990 Model atmospheres for M dwarfs. PhD thesis, Ruprecht Karls University, Heidelberg, Germany.

4 Allard, F., Hauschildt, P. H., Alexander, D. R. & Starrfield, S. 1997 Model atmospheres of very low mass stars and brown dwarfs. *Annu. Rev. Astron. Astrophys.* **35**, 137–177. (doi:10.1146/annurev.astro.35.1.137)
5 Hauschildt, P. H., Allard, F. & Baron, E. 1999 The NextGen model atmosphere grid for $3000 \leq T_{\text{eff}} \leq 10\,000\,\text{K}$. Astrophys. J. 512, 377. (doi:10.1086/306745)
6 Hauschildt, P. H., Allard, F., Ferguson, J., Baron, E. & Alexander, D. R. 1999 The NEXTGEN model atmosphere grid. II. Spherically symmetric model atmospheres for giant stars with effective temperatures between 3000 and 6800 K. Astrophys. J. 525, 871. (doi:10.1086/307954)
7 Seelmann, A. M., Hauschildt, P. H. & Baron, E. 2010 A 3D radiative transfer framework. Astron. Astrophys. 522, 1–4. (doi:10.1051/0004-6361/201014278)
8 Freytag, B., Steffen, M., Ludwig, H.-G., Wedemeyer-Böhm, S., Schaffenberger, W. & Steiner, O. 2011 Simulations of stellar convection with CO5BOLD. J. Comput. Phys. 231, 919–959. (doi:10.1016/j.jcp.2011.09.026)
9 Mould, J. R. 1975 A study of M dwarfs. I. Preliminary model atmospheres. Astron. Astrophys. 38, 283–288. See http://adsabs.harvard.edu/abs/1975A%26A....38..283M.
10 Kui, R. 1991 Model atmospheres for M-dwarfs. PhD thesis, Australian National University, Canberra, Australia.
11 Kivel, B., Mayer, H. & Bethe, H. 1952 Radiation from hot air. I. Theory of nitric oxide absorption. Ann. Phys. 2, 57–80. (doi:10.1006/0003-4916(57)90035-0)
12 Golden, S. A. 1967 Approximate spectral absorption coefficients of electronic transitions in diatomic molecules. J. Quant. Spectrosc. Radiat. Transf. 7, 225–250. (doi:10.1016/0022-4073(67)90067-2)
13 Allard, F. & Hauschildt, P. H. 1995 Model atmospheres for M (sub)dwarf stars. 1: The base model grid. Astrophys. J. 445, 433–450. (doi:10.1086/175708)
14 Ludwig, C. B. 1971 Measurements of the curves-of-growth of hot water vapor. Appl. Opt. 10, 1057–1073. (doi:10.1364/AO.10.001057)
15 Jørgensen, U. G., Jensen, P., Sørensen, G. O. & Aringer, B. 2001 H2O in stellar atmospheres. Astron. Astrophys. 372, 249–259. (doi:10.1051/0004-6361:20010285)
16 Allard, F., Hauschildt, P. H., Miller, S. & Tennyson, J. 1994 The influence of H2O line blanketing on the spectra of cool dwarf stars. Astrophys. J. Lett. 426, L39–L41. (doi:10.1086/187334)
17 Schryber, J. H., Miller, S. & Tennyson, J. 1995 Computed infrared absorption properties of hot water vapour. J. Quant. Spectrosc. Radiat. Transf. 53, 373–380. (doi:10.1016/0022-4073(95)90013-6)
18 Allard, F., Hauschildt, P. H. & Schwenke, D. 2000 TiO and H2O absorption lines in cool stellar atmospheres. Astrophys. J. 540, 1005. (doi:10.1086/309366)
19 Allard, F., Hauschildt, P. H., Alexander, D. R., Tamanai, A. & Schweitzer, A. 2001 The limiting effects of dust in brown dwarf model atmospheres. Astrophys. J. 556, 357. (doi:10.1086/321547)
20 Partridge, H. & Schwenke, D. W. 1997 The determination of an accurate isotope dependent potential energy surface for water from extensive ab initio calculations and experimental data. J. Chem. Phys. 106, 4618–4639. (doi:10.1063/1.473987)
21 Asplund, M., Grevesse, N., Sauval, A. J. & Scott, P. 2009 The chemical composition of the sun. Annu. Rev. Astron. Astrophys. 47, 481–522. (doi:10.1146/annurev.astro.46.060407.145222)
22 Caffau, E., Ludwig, H.-G., Steffen, M., Freytag, B. & Bonifacio, P. 2011 Solar chemical abundances determined with a CO5BOLD 3D model atmosphere. Solar Phys. 268, 255–269. (doi:10.1007/s11207-010-9541-4)
23 Grevesse, N., Noels, A. & Sauval, A. J. 1993 A revision of the solar abundance of dysprosium. Astron. Astrophys. 271, 587. See http://adsabs.harvard.edu/abs/1993A%26A...271..587G.
24 Leinert, C., Haas, M., Allard, F., Wehre, R., McCarthy, D. W., Juhreiss, H. & Perrier, C. 1990 The nearby binary Gliese 866 A/B: orbit, masses, temperature, and composition. Astron. Astrophys. 236, 399–408. See http://adsabs.harvard.edu/abs/1990A%26A...236..399L.
25 Deloffe, X., Forveille, T., Udry, S., Beuzit, J.-L., Mayor, M. & Perrier, C. 1999 Accurate masses of very low mass stars. II. The very low mass triple system GL 866. Astron. Astrophys. 350, L39. See http://adsabs.harvard.edu/abs/1999A%26A...350L..39D.
26 Rayner, J. T., Cushing, M. C. & Vacca, W. D. 2009 The infrared telescope facility (IRTF) spectral library: cool stars. Astrophys. J. Suppl. 185, 289. (doi:10.1088/0067-0049/185/2/289)
27 Casagrande, L., Flynn, C. & Bessell, M. 2008 M dwarfs: effective temperatures, radii and metallicities. Mon. Not. R. Astron. Soc. 389, 585–607. (doi:10.1111/j.1365-2966.2008.13573.x)
28 Golimowski, D. A. et al. 2004 L' and M' photometry of ultracool dwarfs. *Astron. J.* **127**, 3516. (doi:10.1086/420709)
29 Vrba, F. J. et al. 2004 Preliminary parallaxes of 40L and T dwarfs from the US Naval Observatory infrared astrometry program. *Astron. J.* **127**, 2948. (doi:10.1086/383554)
30 Baraffe, I., Chabrier, G., Allard, F. & Hauschildt, P. H. 1997 Evolutionary models for metal-poor low-mass stars. Lower main sequence of globular clusters and halo field stars. *Astron. Astrophys.* **327**, 1054–1069. See http://aa.springer.de/papers/7327003/2301054.pdf.
31 Baraffe, I., Chabrier, G., Allard, F. & Hauschildt, P. H. 1998 Evolutionary models for solar metallicity low-mass stars: mass–magnitude relationships and color–magnitude diagrams. *Astron. Astrophys.* **337**, 403–412. See http://aa.springer.de/papers/8337002/2300403.pdf.
32 Langhoff, S. R. 1997 Theoretical study of the spectroscopy of TiO. *Astrophys. J.* **481**, 1007. (doi:10.1086/304077)
33 Chabrier, G. & Baraffe, I. 1997 Structure and evolution of low-mass stars. *Astron. Astrophys.* **327**, 1039–1053. See http://aa.springer.de/papers/7327003/2301039.pdf.
34 Tsuji, T., Ohnaka, K. & Aoki, W. 1996 Dust formation in stellar photospheres: a case of very low mass stars and a possible resolution on the effective temperature scale of M dwarfs. *Astron. Astrophys.* **305**, L1–L4. See http://adsabs.harvard.edu/abs/1996A%26A...305L...1T.
35 Lodders, K. & Fegley Jr, B. 2006 Chemistry of low mass substellar objects. In *Astrophysics update 2: Topical and timely reviews on astrophysics* (ed. J. W. Mason), pp. 1–28. Chichester, UK: Springer Praxis.
36 Tsuji, T. 2002 Dust in the photospheric environment: unified cloudy models of M, L, and T dwarfs. *Astrophys. J.* **575**, 264. (doi:10.1086/341262)
37 Ackerman, A. S. & Marley, M. S. 2001 Precipitating condensation clouds in substellar atmospheres. *Astrophys. J.* **556**, 872. (doi:10.1086/321540)
38 Burrows, A., Sudarsky, D. & Hubeny, I. 2006 L and T dwarf models and the L to T transition. *Astrophys. J.* **640**, 1063. (doi:10.1086/500293)
39 Lewis, J. S. 1969 The clouds of Jupiter and the NH3H2O and NH3H2S systems. *Icarus* **10**, 365–378. (doi:10.1016/0019-1035(69)90091-8)
40 Rossow, W. B. 1978 Cloud microphysics: analysis of the clouds of Earth, Venus, Mars and Jupiter. *Icarus* **36**, 1–50. (doi:10.1016/0019-1035(78)90072-6)
41 Lunine, J. I., Hubbard, W. B., Burrows, A., Wang, Y.-P. & Garlow, K. 1989 The effect of gas and grain opacity on the cooling of brown dwarfs. *Astrophys. J.* **338**, 314–337. (doi:10.1086/167201)
42 Helling, C., et al. 2008 A comparison of chemistry and dust cloud formation in ultracool dwarf model atmospheres. *Mon. Not. R. Astron. Soc.* **391**, 1854–1873. (doi:10.1111/j.1365-2966.2008.13991.x)
43 Tsuji, T., Nakajima, T. & Yanagisawa, K. 2004 Dust in the photospheric environment. II. Effect on the near-infrared spectra of L and T dwarfs. *Astrophys. J.* **607**, 511. (doi:10.1086/383300)
44 Allard, F., Hauschildt, P. H., Alexander, D. R., Tamanai, A. & Schweitzer, A. 2003 Model atmospheres and spectra: the role of dust. In *Brown dwarfs* (ed. E. Martín). IAU Symp. Ser., no. 211, p. 325. San Francisco, CA: Astronomical Society of the Pacific.
45 Saumon, D. & Marley, M. S. 2008 The evolution of L and T dwarfs in color–magnitude diagrams. *Astrophys. J.* **689**, 1327. (doi:10.1086/592734)
46 Helling, C., Dehn, M., Woitke, P. & Hauschildt, P. H. 2008 Consistent simulations of substellar atmospheres and nonequilibrium dust cloud formation. *Astrophys. J. Lett.* **675**, L105. (doi:10.1086/533462)
47 Witte, S., Helling, C., Barman, T., Heidrich, N. & Hauschildt, P. H. 2011 Dust in brown dwarfs and extra-solar planets. *Astron. Astrophys.* **529**, A44. (doi:10.1051/0004-6361/201014105)
48 Freytag, B., Allard, F., Ludwig, H., Homeier, D. & Steffen, M. 2010 The role of convection, overshoot, and gravity waves for the transport of dust in M dwarf and brown dwarf atmospheres. *Astron. Astrophys.* **513**, A19. (doi:10.1051/0004-6361/200913354)
49 Noll, K. S., Geballe, T. R. & Marley, M. S. 1997 Detection of abundant carbon monoxide in the brown dwarf Gliese 229B. *Astrophys. J. Lett.* **489**, L87. (doi:10.1086/310954)
50 Griffith, C. A. & Yelle, R. V. 1999 Disequilibrium chemistry in a brown dwarf’s atmosphere: carbon monoxide in Gliese 229B. *Astrophys. J.* **519**, L85–L88. (doi:10.1086/312103)
51 Tsuji, T., Yamamura, I. & Sorahana, S. 2011 Akari observations of brown dwarfs. II. CO₂ as probe of carbon and oxygen abundances in brown dwarfs. Astrophys. J. 734, 73. (doi:10.1088/0004-637X/734/2/73)

52 Saumon, D., Marley, M. S., Cushing, M. C., Leggett, S. K., Roellig, T. L., Lodders, K. & Freedman, R. S. 2006 Ammonia as a tracer of chemical equilibrium in the T7.5 dwarf Gliese 570D. Astrophys. J. 647, 552. (doi:10.1086/505419)

53 Barman, T. S., Hauschildt, P. H. & Allard, F. 2001 Irradiated planets. Astrophys. J. 556, 885. (doi:10.1086/321610)

54 Barber, R. J., Tennyson, J., Harris, G. J. & Tolchenov, R. N. 2006 A high-accuracy computed water line list. Mon. Not. R. Astron. Soc. 368, 1087. (doi:10.1111/j.1365-2966.2006.10184.x)

55 Homeier, D., Hauschildt, P. H. & Allard, F. 2003 Methane opacities in T dwarf atmospheres. In Stellar atmosphere modeling (eds I. Hubeny, D. Mihalas & K. Werner). ASP Conf. Ser., vol. 288, p. 357. San Francisco, CA: Astronomical Society of the Pacific.

56 Sharp, C. M. & Burrows, A. 2007 Atomic and molecular opacities for brown dwarf and giant planet atmospheres. Astrophys. J. Suppl. 168, 140. (doi:10.1086/508708)

57 Tashkun, S. A., Perevalov, V. I., Teffo, J.-L., Bykov, A. D., Lavrentieva, N. N. & Babikov, Y. L. 2004 CDSD-1000: the high-temperature carbon dioxide spectroscopic databank and information system. In Proc. 14th Symp. on High-Resolution Molecular Spectroscopy (eds. L. N. Sinitsa & S. N. Mikhailenko). Proc. SPIE 5311, 102–113. (doi:10.1117/12.545199)