The low-mass stellar population in the young cluster Tr 37

Disk evolution, accretion, and environment

Aurora Sicilia-Aguilar1, Jinyoung Serena Kim2, Andrej Sobolev3, Konstantin Getman4, Thomas Henning5, and Min Fang1

1 Departamento de Física Teórica, Facultad de Ciencias, Universidad Autónoma de Madrid, 28049 Cantoblanco, Madrid, Spain
2 Steward Observatory, University of Arizona, 933 North Cherry Avenue, Tucson, AZ 85721-0065, USA
3 Astronomical Observatory, Ural Federal University, Lenin Avenue 51, 620000 Ekaterinburg, Russia
4 Department of Astronomy & Astrophysics, 525 Davey Laboratory, Pennsylvania State University, University Park PA 16802, USA
5 Max-Planck-Institut für Astronomie, Königstuhl 17, 69117 Heidelberg, Germany

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ABSTRACT

Aims. We present a study of accretion and protoplanetary disks around M-type stars in the 4 Myr-old cluster Tr 37. With a well-studied solar-type population, Tr 37 is a benchmark for disk evolution.

Methods. We used low-resolution spectroscopy to identify and classify 141 members (78 new ones) and 64 probable members, mostly M-type stars. Ho emission provides information about accretion. Optical, 2MASS, Spitzer, and WISE data are used to trace the spectral energy distributions (SEDs) and search for disks. We construct radiative transfer models to explore the structures of full-disks, pre-transition, transition, and dust-depleted disks.

Results. Including the new members and the known solar-type stars, we confirm that a substantial fraction (~2/5) of disks show signs of evolution, either as radial dust evolution (transition/pre-transition disks) or as a more global evolution (with low small-dust masses, dust settling, and weak/absent accretion signatures). Accretion is strongly dependent on the SED type. About half of the transition objects are consistent with no accretion, and dust-depleted disks have weak (or undetectable) accretion signatures, especially among M-type stars.

Conclusions. The analysis of accretion and disk structure suggests a parallel evolution of dust and gas. We find several distinct classes of evolved disks, based on SED type and accretion status, pointing to different disk dispersal mechanisms and probably different evolutionary paths. Dust depletion and opening of inner holes appear to be independent processes: most transition disks are not dust-depleted, and most dust-depleted disks do not require inner holes. The differences in disk structure between M-type and solar-type stars in Tr 37 (4 Myr old) are not as remarkable as in the young, sparse, Coronet cluster (1–2 Myr old), suggesting that other factors, like the environment/interactions in each cluster, are likely to play an important role in the disk evolution and dispersal.

Finally, we also find some evidence of clumpy star formation or mini-clusters within Tr 37.

Key words. stars: pre-main sequence – protoplanetary disks – stars: late-type

1. Introduction

The evolution of accretion disks around low-mass stars, and the way protoplanetary disks cease to accrete and disappear, presumably after forming planets, is since long matter of discussion. Protoplanetary disks around solar-type stars have typical lifetimes of few Myr (Haisch et al. 2001; Sicilia-Aguilar et al. 2006a; Hernandez et al. 2007; Fedele et al. 2010), but the high variety of disk morphologies at a given age suggests that disk evolution is controlled by several mechanisms. Various physical processes can be invoked for disk removal, including grain growth, photoevaporation, and planet formation. In addition, other external parameters may also contribute to define the path followed by a dispersing disk, such as age, stellar mass, stellar and planetary companions, initial conditions, cluster environment, crowdedness in the star-forming region, and angular momentum of the collapsing core (Hartmann et al. 2006; Bouwman et al. 2006; Alexander & Armitage 2009; Fang et al. 2013a; Sicilia-Aguilar et al. 2013; Dullemond et al. 2006). Further processes (ejection in multiple systems, photo-erosion of cores by massive nearby stars; Bate et al. 2005, 2012; Whitworth & Zinnecker 2004) may also play a role in the formation of low-mass systems. The unexpected and not yet understood relation between stellar mass and accretion rate (dM/dt ∼ M2; Natta et al. 2004; Hartmann et al. 2006; Fang et al. 2009; but also dM/dt ∼ M1.5; Barentsen et al. 2011) is consistent with these multiple parameters affecting disk evolution (Clarke & Pringle 2006; Gatti et al. 2008; Herczeg & Hillenbrand 2008).

Multimwavelength observations of disks around faint objects are a challenge for current instrumentation. There is a lack...
of combined optical photometry, spectroscopy, and complete IR data for large samples of very low-mass objects. The first observations of disks around very low-mass stars and brown dwarfs (BD) suggested that they are lower-mass analogs of the typical T Tauri disks. Their disks are flared, show active accretion with strong Hα emission, silicate emission from small grains in the disk atmosphere, processed and crystalline silicates, and dust continuum emission down to the far-IR and millimeter waves (Muench et al. 2001; Klein et al. 2003; Mohanty et al. 2004; Jayawardhana et al. 2005; Apai et al. 2005; Scholz et al. 2007; Scholz & Jayawardhana 2008; Harvey et al. 2010, 2012a,b). Many of these observations were biased towards the most luminous disks and stronger accretors. Spitzer data suggested that there is an important fraction of harder-to-detect, settled, low-mass, and transitional disks with inner gaps around the very low-mass objects (masses <0.2 M☉, down to the BD regime; Morrow et al. 2008). IR silicate spectroscopy of M-type stars also suggested differences in innermost disk evolution (Kessler-Silacci et al. 2007; Sicilia-Aguilar et al. 2007, 2008; Pascucci et al. 2009). Differences in the disk structure, dead zones, and accretion mechanisms for the lower-mass objects could also change the evolution of the disk and the formation of planets around very low-mass stars (Hartmann et al. 2006).

More recent observations suggest that the stellar mass is not the only parameter that controls the disk structure and subsequent evolution. Studies of stars in sparse clusters vs. more populous star-forming regions point to differences in the disk fraction vs. age trend. Sparse clusters would have relatively lower disk fractions at an early age, but these disks may survive for longer timescales compared to more massive regions (Fang et al. 2013a). The Herschel Space Telescope has traced the typical sizes and structures where young stars are formed (Arzoumanian et al. 2011; Hacar et al. 2013), revealing the details of star-forming filaments. Surprisingly, some sparse associations, instead of being quiet low-mass star-forming regions, appear to be very crowded, active, and interactive already at a very early stage (Sicilia-Aguilar et al. 2013). The cluster dynamics, interactions, and angular momentum in the collapsing cloud could also affect the initial mass function (IMF) and the disk properties and subsequent evolution (Hsu et al. 2012, 2013; Becker et al. 2013; Dullemond et al. 2006).

With this work we want to address disk evolution and its dependency on stellar mass/spectral type and environment by studying the low-mass members in the Tr 37 cluster. Tr 37 is located at 870 pc distance (Contreras et al. 2002) and part of the Cep OB2 complex (Platais et al. 1998; Patel et al. 1995, Patel). The cluster is a key region for disk evolution studies due to its intermediate age (∼4 Myr; Sicilia-Aguilar et al. 2005) compared to the typical disk lifetimes (3–10 Myr; Sicilia-Aguilar et al. 2006a; Hernández et al. 2007; Fedele et al. 2010). Multiwavelength studies (Sicilia-Aguilar et al. 2004, 2005, 2006a,b; from now on SA04, SA05, SA06a, and SA06b) have targeted the solar-type population in the region, finding evidence of substantial disk evolution and dust processing (SA06a; Sicilia-Aguilar et al. 2007, 2011b, from now on SA11). Detailed Hα photometry surveys (Barentsen et al. 2011) revealed an extended population of accreting stars, some of which could be younger than the main cluster (SA05, Getman et al. 2012).

The disk structure for the solar-type population has been already extensively discussed in SA06a and SA11. Using a combination of optical (photometry and spectroscopy), Spitzer (IRAC, MIPS, and IRS), and millimeter-wave (IRAM) data, together with simple radiative transfer models (RADMC; Dullemond & Dominik 2004), we constrained disk parameters and deviations from typical, uniform, flared disks. Here we present a detailed multwavelength study of the low-mass population in Tr 37. By using optical spectroscopy, we were able to classify more than 200 objects among cluster members and probable members, including 78 newly identified members, most of them M-type stars. Combining this information with our previous optical photometry and Spitzer IRAC and MIPS data, we analyze the disk characteristics of the objects and put them in context comparing them to our previous study of the solar-type stars. All observations are presented in Sect. 2. In Sect. 3 we examine the membership of the candidates and derive their fundamental stellar properties. In Sect. 4 we explore the implications of the newly discovered objects for disk dispersal and evolution. Finally, Sect. 5 summarizes our results.

2. Observations and data reduction

2.1. Optical and IR data and candidate selection for spectroscopy

This study aimed to complete the previous work of SA05/SA06a/SA11 on the disks around solar-type stars in the Tr 37 cluster, which was approximately complete for stars with spectral types K4–M2, by addressing the M-type population in the cloud. Optical photometry and spectroscopy are required to confirm the cluster membership, to obtain spectral types, to detect accretion, and to estimate extinction, age, and stellar mass. The starting point of the target selection was the deep optical photometry obtained at Calar Alto Observatory with the LAICA camera on the 3.5 m telescope, consisting of deep observations with the standard UVRI Johnsons filters (see Sicilia-Aguilar et al. 2010 for a detailed description of the observations and data reduction). The data had a large dynamical range, being complete in the range U ∼ 15–21, V ∼ 13–21, R1 ∼ 12–21, and I1 ∼ 11–20 mag. The low-mass candidates relevant for our survey were in general too faint to be detected with the U band filter. The requirement for the objects to be preferably candidate M-type members resulted in a selection of objects with R1 = 17–20.5 mag, considering that the extinction over Tr 37 is moderate and relatively uniform (AV = 1.56 ± 0.55 mag). Examining the 2MASS (Cutri et al. 2003) counterparts of our optical targets allowed to refine the selection of objects consistent with diskless late-type stars and the locus of classical T Tauri stars (CTTS) in the J – H vs. H – K diagram (Bessell & Brett 1988; Meyer et al. 1997), although for M-type stars with evolved disks, the excesses in H and K bands are usually very small or negligible. The 2MASS coordinates were also used for the later spectroscopy, given the strong requirements of multiobject spectrographs.

To complete the spectral energy distributions (SEDs) of the candidates, we re-reduced the Spitzer IRAC and MIPS observations (from both our previous datasets in SA06a, and Spitzer archival data), following the method described in SA11. These new photometry is not significantly different from that of SA06a, although the use of a more recent pipeline with improved flat fielding, and smaller apertures result in more accurate magnitudes for the in-cloud sources and a better detection of very faint objects in the MIPS 24 µm maps. To ensure that contamination by cloud emission or ghosts remains minimal, we visually inspected all the candidates. In particular, detailed inspection of the 24 µm data was necessary to remove objects suffering from nebular emission.
For the spectroscopic followup, we selected the sources with ages under 100 Myr that were consistent with solar- and M-type stars at 870 pc distance according to the $V$ vs. $V-I$ and $V-R$ diagrams and the Siess et al. (2000) isochrones. This produced a list of about 400 targets, of which approximately 100 had IR excesses consistent with circumstellar disks and that lacked spectroscopic characterization. We assigned priority to allocating the fibers and slits to objects with IR excess, followed by those with isochrone ages $<10$ Myr, and finally those between $10-100$ Myr. All the observed targets are listed in Table A.1, together with the relevant information about Hα, Li I emission, presence of disks, spectral type, extinction, and membership (see Sect. 3 for details). A few remaining fibers were assigned to previously known objects, or to objects with strong IR excesses but lacking optical data. This study is thus biased towards objects with disks, so we cannot estimate the disk fraction. The main advantage is that we covered more than 90% of the candidate M-type stars with IR excesses, obtaining a superb dataset to explore the various disk structures. The photometry data (optical, 2MASS, and Spitzer) are listed only for the members and probable members in Tables A.2 and A.3.

For completeness, after the survey we also checked the WISE database for counterparts to our objects. The Spitzer data are preferred because the larger WISE beam often suffers from strong contamination due to cloud emission, resulting in overestimated fluxes, especially in bands 3 and 4. Nevertheless, band 3 ($12 \mu m$), together with IRAC 8 $\mu m$, offers valuable information regarding the silicate emission. Whenever there was a good agreement between IRAC and WISE bands 1 and 2, MIPS 24 $\mu m$ and WISE band 4 ($22 \mu m$), we assume there is no significant contamination and thus use the 12 $\mu m$ WISE data. The WISE fluxes were obtained applying the Wright et al. (2010) color corrections depending on the Spitzer SED shape. The WISE data used in this project are also listed in Table A.4.

The spectroscopic properties of the observed objects and membership criteria are described in Sect. 3. Thanks to the strong candidate selection criteria, the efficiency of the survey was very high, with about 90% of candidates with isochronal ages younger than 10 Myr and IR excesses being confirmed as cluster members. The success rate fell down to ~5-20% among the objects without excess. A number of objects with very low signal-to-noise ratio (S/N) could not be ruled out nor confirmed as members (see Table A.1). A few examples of SEDs of members with IR excesses are shown in Fig. 1. The SEDs of all members and probable members with IR excesses are displayed in Figs. B.1–B.6, and the SEDs classification criteria are discussed in Sect. 4.1.

2.2. SCORPIO multislit/6 m Russian SAO RAS Telescope

One observational campaign took place with the Spectral Camera with Optical Reducer for Photometric and Interferometric Observations (SCORPIO) multislit spectrograph (Afanasiev & Moiseev 2005) at the 6 m telescope of the Special Astrophysical Observatory of the Russian Academy of Sciences, in the Northern Caucasus. The observations were obtained during 2008 October 4–6. SCORPIO is a multislit spectrograph with 16 movable long slits (each of 18′′) that can be placed over a 2.9 × 5.9 arcmin$^2$ field. We observed 3 configurations, with slits allocated to cover between 12 and 16 candidate objects. We used the grism VPHG1200r, with a spectral coverage 5700–7500 Å and a 5 Å/pix resolution. Due to the poor weather conditions, we had to reduce the exposure time to 3 × 15–10 min per field.

The observations were reduced following standard IRAF1 procedures within the *spec2d* package, in particular, routines within the *tpospec* and *apextract* packages. The spectra were bias-corrected, flat fielded, and extracted. The wavelength calibration was performed with a Ne lamp, observed with the same slit configuration as the datasets. The sky subtraction was done by using the sky spectra adjacent to the star in the slit. A total of 18 new member candidates were identified among the observed objects. The main limitation of these observations was the poor S/N of the data, due to the poor weather conditions, that did not allow us to clearly detect the L1 absorption in most of the candidates. We therefore re-observed many of these objects in our subsequent spectroscopic program with the MMT. The advantage of using

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1. IRAF is distributed by the National Optical Astronomy Observatories, which are operated by the Association of Universities for Research in Astronomy, Inc., under cooperative agreement with the National Science Foundation.
Fig. 2. Some examples of spectra taken with the MMT/Hectospec. The Hα features are cut to reveal more details of the whole spectra. Note that the spectra still show some telluric lines in emission (mostly at 5577 and 5004 Å). The flux is in arbitrary instrumental units.

### 2.3. Hectospec/MMT spectroscopy

The candidate objects were observed during two campaigns with the multifiber spectrograph Hectospec on the 6.5 m MMT telescope in Mount Hopkins, AZ. The first set of observations were taken on 2009 July 14 and 19, and the second one on 2010 May 17 and 19. The weather conditions were fair during both campaigns. Hectospec is a multifiber spectrograph with 300 fibers that can be allocated to individual objects or sky over a 1 degree field of view. We took a total of 3 configurations, each one including between 180 and 220 candidate objects, plus numerous sky positions. We used the 270 gpm grating, with a spectral coverage from ∼3650–9200 Å and a typical resolution of 5 Å. This is the same setup we used for the observations of the solar-type stars (SA05). For each field, we took 3 × 30 min exposures, followed by sky observations obtained by shifting the whole configuration by 5″ and exposing 3 × 10 min. These sky observations are particularly useful to attain a good sky subtraction in the cases of faint targets or objects in regions with complicated background, since they provided the sky spectrum in a region very close to the objects, and obtained with the same fiber.

The bias, flat fielding, and wavelength calibration were done with standard IRAF tasks within the dofibers package. For the wavelength calibration, we used a HeNeAr calibration lamp. Due to the offsets between the even- and odd-numbered fibers, the wavelength calibration was done separately for the even and odd fibers. The one-dimensional spectra were extracted and combined before the sky subtraction. The sky subtraction can be a problem in regions like Tr 37, where substantial, spatially-variable emission from the H II region can affect the Hα line, one of our main membership diagnostics. For each night, we classified all the sky spectra in “bright”, “medium”, and “faint”, according to the intensities of the nebular emission lines (especially, Hα, [N II], and [O I]), combining the three sets to create bright, medium, and faint sky templates. Each template resulted from a minimum of 20 spectra. The three templates were subtracted from all the objects, and we then examined the success of each sky subtraction by checking the night sky and nebular emission lines. We then selected the best-subtracted spectrum. For about a fourth of the objects, none of the three sky subtractions provided a good result (either by excess or by defect). In those cases, we used the wavelength-calibrated individual sky spectra taken in the proximity of the star, which improved the subtraction of the nebular and night sky lines, although the difference in exposure time resulted in a higher noise than in the cases where the sky template spectra were applied. Some examples of spectra are shown in Fig. 2.

### 2.4. CAFOS/2.2 m Calar Alto narrow band imaging

Since our previous spectroscopic surveys had revealed substantial forbidden line emission in the Tr 37 globule, we also performed a narrow band imaging in the [S II] lines using the Fabry-Pérot interferometer with the CAFOS wide-field imager mounted on the 2.2 m telescope in Calar Alto. The data were obtained as part of a Director Discretionary Time program on 28 August 2009.

We obtained images centered in the Tr 37 globule, around the CTTS 14−141, RA(2000) Dec(2000) 21:36:49.41 +57:31:22.0, which showed remarkable variable forbidden line emission in our previous spectroscopic surveys. We obtained 3 dithered 600 s exposures with the Fabry-Pérot interferometer at 6716, 6730, and 6750 Å, resulting in narrow-band images centered on the two [S II] lines and a continuum observation with the same width in the line-free region around 6750 Å. An additional 3 × 10 s Johnsons R band image was also obtained for comparison. The data were reduced (bias, flat field) and combined according to the standard procedures using IRAF packages. Finally, a line-only image was obtained by subtracting the continuum 6750 Å
image from the [S II] ones. To improve the S/N, we combined both [S II] lines for the final result. We did not attempt any flux calibration, which is difficult due to the properties of the filter and to the fact that there is a wavelength drift over the CAPOS field. The resulting data clearly reveals several shocks in the field, not only associated to 14–141, but also to several of the protostars and embedded objects in the cloud (see Fig. 3).

3. Analysis

3.1. Spectral types and extinction

As a preliminary step to determine the membership and properties, we obtained spectral types for all the observed objects, which are listed in Table A.1. The spectral types of the objects with good S/N were derived using the combinations of indices employed for the classification of K-M3 stars in Cep OB2 by Sicilia-Aguilar et al. (2004), and a similar scheme for M3–M8 stars as it was used for classification of the Coronet cluster members (Sicilia-Aguilar et al. 2008). The classification of M3–M8 stars was based on the indices defined by Kirkpatrick et al. (1995), Martín et al. (1996), and Riddick et al. (2007).

From the indices therein, we selected the best ones that did not suffer from strong problems due to atmospheric features, extinction, or lack of S/N in the blue (see Table 1). The calibration in Sicilia-Aguilar et al. (2004) is done in effective temperature (see Table 2). The effective temperatures ($T_{\text{eff}}$, in units of $10^3$ K) are derived as $T_{\text{eff}} = T_0 + b \times (I - I_0)$, where $I$ is the measured index, and $T_0$, $I_0$, and $b$ are the zero-point values of effective temperature and index, and the slope of the relation resulting from the fit of several standard stars (Sicilia-Aguilar et al. 2004). The index $I$ is obtained as a function of the flux in different wavelength ranges ($F_{\lambda_1-\lambda_2}$) is the flux between the wavelengths $\lambda_1$ and $\lambda_2$ in Å). The values of $T_{\text{eff}}$ are transformed to spectral types using the calibration by Kenyon & Hartmann (1995). The spectral types of objects observed with SCORPIO were derived by direct comparison to standard stars, given their reduced wavelength coverage. We also note that since the SCORPIO spectra do not include the main features for classification of G/K stars, stars other than M-type observed by SCORPIO have very uncertain spectral types.

The indices are not applicable out of a given spectral range. Therefore, we did a preliminary visual classification of the stars as “earlier than G”, “GK-type”, and “M-type” before computing the spectral types. Due to our selection criteria, most of the stars found to be earlier than $G$ are unlikely members, and they are simply listed in Table A.1 as “early type” (E). For the objects classified as “GK-type”, we used the MgI and CaI indices. The number of indices applicable to G-type and early-K stars is also very low, so the spectral types of objects earlier than K5 were refined by comparison to standard stars. For the objects classified as “M-type”, we used first the TiO 6185 index to determine whether the object was earlier than or later than M3. For objects earlier than M3, we then used the CaI and TiO71 indices to refine the spectral type. For M3 or later stars, we used the PC1, R1, and R3 indices to derive a preliminary spectral type, which was refined by using the TiO 8465, VOA, and VOB, and PC2 indices, if within their applicable range. For each star, the final spectral type is calculated as the average of the types obtained with the applicable indices. In all cases, we performed a visual comparison with standard stars to check that the estimated spectral type is in agreement with the appearance of the spectrum. The few cases where there was a discrepancy between the derived type and the visual inspection could be tracked to different
problems (presence of stellar or sky emission lines, lack of S/N in part of the spectrum, contamination by scattered light) that were addressed by removing the discrepant indices, and using the noncontaminated ones to derive the final spectral type.

In case of the objects with more than one spectra, we derived the spectral type for each one and consider the result of the best S/N spectrum (if the quality of both was very different). There is a very good agreement between all estimates (within 1–2 subtypes, depending on spectral type and data quality). Spectral types had been previously derived for 14 among the 17 objects in common with our previous campaigns. Among these, we recover the previous spectral type (±1 subtype) in 9 cases, 2 cases are off by up to 2 subtypes, and 3 are off by up to 3 subtypes (corresponding mostly to faint M-type stars with low S/N in this or in previous surveys and/or limited spectral coverage). This confirms our typical estimated error of 1–2 subtypes, strongly depending on the S/N. In some cases, real spectroscopic variations were observed between 2009 and 2010, probably due to variations in the accretion activity and in the stellar spots. Two examples of this behavior are 213701319+573418289 (M3.0/M1.0, 2009 spectrum later than the 2010 one) and 213710877+573846877 (M4.5/M2.0, later in 2009).

We obtained the extinction of each object by comparing the observed optical data at VRI with the theoretical colors for young Taurus stars with the same spectral type (Kenyon & Hartmann 1995), and applying the relations of Cardelli et al. (1989) to derive $A_V$ from $E(V - I)$ and $E(V - R)$ for the Johnsons filters. Whenever VRI photometry was available, we used both $E(V - I) = 0.521 A_V$ and $E(V - R) = 0.249 A_V$, estimating $A_V$ as the average, with an error corresponding to the standard deviation. In the cases where one optical band was missing, we used the available data ($V - I$, $V - R$, or $R - I$) to compute the extinction, and estimated the error considering the typical spectral typing error. In the cases where no optical data (or only one optical band) were available, the extinction was derived from the 2MASS photometry, considering the theoretical colors from Bessell & Brett (1988) and the corresponding relations for $E(J - H) = 0.092 A_V$ and $E(H - K) = 0.076 A_V$. For the few objects that are confirmed to be members but have no reliable spectral types, we take the average cluster extinction ($A_V = 1.6$ mag; SA05). The final extinctions are listed in Table A.1.

### 3.2. Membership

The membership of the stars was established by a combination of several indicators. Typical youth indicators are the Li I absorption at 6708 Å and the Hα equivalent width (EW). The EW of both Hα and Li I are listed in Table A.1. Table A.5 lists other lines that have been observed in some members, mostly strongly accreting CTTS where the whole Balmer series is visible, together with the Ca II IR triplet, He I lines, Ca II H and K lines, and some Fe I and O I emission (Hamann & Persson 1992; Appenzeller et al. 2005), plus a few objects with forbidden [S II] and [O I] lines related to shocks (Hamann 1994).
most reliable indicator of youth for late-type stars is the Li I absorption at 6708 Å (Fig. 4). Objects with clear Li I detection are labeled as sure members, and objects with good S/N at 6708 Å and no evidence of Li I absorption as labeled as nonmembers. Unfortunately, the Li I line is not always detectable in these faint objects, so we need additional membership criteria. The Hα line in emission is the main characteristic of T Tauri stars. In general, we followed the criteria of White & Basri (2003) regarding accretion and the distinction between classical and weak-line T Tauri stars (CTTS/WTTS). For the objects observed with the multifiber spectrograph Hectospec, sky subtraction is complicated due to the strong nebular emission in Tr 37, so low Hα EW values are uncertain. Young, nonaccreting M-type stars can be hard to classify based on their Hα emission (Cardelli et al. 1989) is also displayed. Errorbars (considering the photometric and $A_V$ errors) are also displayed, although they are often smaller than the symbols.

Finally, since this survey was fully independent of the X-ray observations of the IC 1396A globule by Getman et al. (2012), we also checked the objects in common with X-ray observations. A total of 23 of the X-ray YSO candidates and 16 of the non-Xray YSO with disks were found in common with our member selection, all labeled as sure and probable members. Four probable members with X-ray detections were thus upgraded to sure members. Three further objects (one probable member, two probably nonmembers) were listed among uncertain X-ray detections, and we are keeping their original membership classification.

Placing the new members and probable members in the $V$ vs. $V-I$ diagram (Fig. 5), we find that their ages are consistent with the mean age of 4 Myr derived previously (SA05). Since the Siess et al. isochrones are given in the Cousins system, we transformed them into the Johnsosn system following the prescriptions of Fernie (1983, for blue stars, $V-I_c < 1.5$ mag), and Getman et al. (2012, for red stars, $V-I_c > 1.5$ mag). A few of the objects appear below the 10 Myr isochrone, but in most cases they are stars with flared disks and high (or even anomalous) extinction being thus probably UXor candidates (all labeled in Table A.1). In addition, early K and G stars tend to deviate from typical isochrone models, probably due birthline uncertainties (Hartmann 2003). Some of the probable members that have anomalous positions in the $V$ vs. $V-I$ diagram may be
contaminants, giving us an idea of the uncertainty in the membership of the objects marked with "P".

Taking into account the five different criteria (Li I absorption, Hα emission, disk excess, extinction, and X-ray), from the initial collection of 565 objects observed, we arrive to a collection of 141 sure members, plus 64 probable members and 33 probable nonmembers. The rest of the 308 observed stars are rejected as nonmembers or marked as uncertain cases on the basis of low S/N (19 in total). A total of 28 identified members are in common with the Hα survey of Barentsen et al. (2011), and 17 objects corresponded to re-observations of previously identified members in SA05 and SA06b. One further object was suggested to be a YSO by Morales-Calderón et al. (2009). Excluding the objects that had been previously suggested to be members by the different authors, we arrive to 78 newly identified members and 64 probable members (to be confirmed with future followup observations), mostly M-type stars and GK members with extinctions higher than the cluster average and with protoplanetary disks.

**4. Discussion**

4.1. Disk structure and the accretion/disk connection

As we had found in our previous papers (SA06a; Sicilia-Aguilar et al. 2007), there is a large variety of disk morphologies in Tr 37. We have classified the new disks following the scheme developed in our Spitzer IRS-based study (SA11). The main criteria for defining the disk evolutionary state are the presence of inside-out evolution, as in transitional and pre-transitional disks, and the evidence of generalized small-dust depletion, as in homologously depleted disks (Currie et al. 2009). A comparison of the disk slope at short and long wavelengths is fundamental to distinguish disks with evidence of inside-out evolution, although the silicate feature at 10 μm also plays an important role. Objects with inner gaps or partially cleared disks, such as pre-transitional disks (Espaillat et al. 2010), have weak near-IR fluxes, strong silicate features, and strong mid-IR excesses, and are especially hard to identify without spectroscopic information. Low fluxes at 24 μm and longer wavelengths are a key to identify small-dust-depleted disks.

To quantify inner disk evolution, we consider \( \alpha (\lambda_1 - \lambda_2) \), defined as the SED slope between two wavelengths:

\[
\alpha (\lambda_1 - \lambda_2) = \frac{\log (\lambda_1 F_{\lambda_1}) - \log (\lambda_2 F_{\lambda_2})}{\log \lambda_1 - \log \lambda_2} \tag{1}
\]

Objects with nearly photospheric colors ([3.6]–[4.5] < 0.2 mag, \( \alpha (3.6-4.5) \sim -3.0 - 2.13 \)) but significant excess at longer wavelengths are classified as transition disks (TD). Disks with moderate to low near-IR excess ([3.6]–[4.5] \sim 0.2 - 0.4 mag, \( \alpha (3.6-4.5) \sim -2.13 - 1.30 \)) and a change in the sign of the slope between 8–12 and 24 μm are labeled as “kink” disks (according to SA06a terminology) and are good candidates for pre-transitional disks (PTD).

Regarding the small-dust mass, and based on our previous results (SA11), we consider objects with low IR fluxes at all wavelengths and reduced 24 μm fluxes as globally dust-depleted disks. Dust-depleted disks have near-IR in the PTDF range or lower, but with similarly steep slopes at all other wavelengths and thus very low 24 μm fluxes (for K- and M-type stars, \( \lambda F_\lambda \lesssim 2 \times 10^{-13} \text{ erg cm}^{-2} \text{s}^{-1} \)). We labeled the objects in Table A.3 according to the disk classification. Figure 6 is color-coded to show the different disk types vs. the SED slopes at different wavelengths. The final classification also requires visual inspection of all the SEDs to detect the few cases that satisfy the mentioned criteria but display SEDs that evidently do not correspond to the presumed disk class. Objects with uncertain photometry/excesses are also excluded from the following analysis and discussion.

Several studies have suggested a strong connection between the IR excesses typical of protoplanetary disks and active accretion (e.g. Sicilia-Aguilar et al. 2006b, 2010; Fedele et al. 2010) while some others suggested that the differences between nonaccreting, WTTS and CTTS in terms of disks are minimal (e.g. Cieza et al. 2007; Oliveira et al. 2013). Differences in accretion between transitional and pre-transitional disks have also been matter of debate (Najita et al. 2007; Muzerolle et al. 2010; Sicilia-Aguilar et al. 2010; Fang et al. 2009, 2013b; Merín et al. 2010). Tr 37 is an interesting region, having a large number of disks with signs of evolution and a disk fraction slightly below 50% (SA06a). A first approach including all known low-mass stars with spectral types G, K, M, suggests a high correlation between the Hα EW, the SED slopes, and our disk classification (Fig. 7; the disk classification is shown in the color-code). Figure 8 displays the Hα EW vs. the effective temperatures (\( T_{\text{eff}} \)) for the whole sample of low-mass members, also color-coded according to their disk types. Instead of separating the objects in CTTS/WTTS according to their Hα EW, we directly considered the measured Hα EW within each disk class, and performed a double-side Kolmogorov-Smirnov (KS) test to find out the probability that the Hα EW of different types of objects are drawn from the same sample.

Table 3 summarizes the results of the KS test. We find unmistakable evidence of a different distribution of Hα EW between objects with disks (including all types of disks) and without disks (no evidence of excess emission at any wavelength), with probabilities nominally 0 that both distributions are drawn from the same sample. We also find significant differences between the Hα EW of TD and full-disks, mostly related to the fact that between 40 and 60% of the TD have Hα values consistent with nonaccreting stars\(^4\), while this is extremely rare among full-disks. Fang et al. (2009, 2013b) found that the fraction of strong accretors (\( \text{Hα EW} > 2 \times \text{Hα}_{\text{threshold}} \) for a given spectral type) in the Orion Lynds 1641 and 1630N clouds is around 60% among full dishes, while only 21–26% of the TD have comparable high accretion. Fang et al. (2013b) inventoried a sample of \( \sim 130 \) YSOs in Lynds 1641. Their sample includes more than 60 WTTS/TD objects which were observed with spectroscopy measuring Hα EW. Among their PTD/TD sample, 36 objects can be classified as TD, and 25 sources as PTDs according to our criteria. The fractions of accretors in both samples are \( 42 \pm 11% \) and \( 52 \pm 14\% \), respectively. Therefore, the results in the Orion Lynds clouds are consistent with Tr 37.

Dust-depleted disks have also significantly lower Hα EW, and thus significantly lower accretion rates (sometimes consistent with no accretion) than normal full-disks. This supports the idea of a strong connection between gas and dust evolution,

\( ^3 \) These include 11–1499, which was labeled as probably nonmember following the criteria in this study, but had Li I absorption in our better S/N spectra from SA05 and X-ray emission according to Getman et al. (2012), so we consider it as member, and the previously identified probable member 21–851, now rejected as cluster member.

\( ^4 \) Note that although Hα EW is a good indicator of accretion, further data like high-resolution Hα spectroscopy or \( \text{U} \) band photometry is needed to rule out cases with very low accretion rates.
where both signs of inside-out evolution (in TD) and global dust depletion appear to be related to lower accretion rates. Observations of evolved disks and accretion signatures in older clusters, like γ Velorum (Hernández et al. 2008; Jeffries et al. 2009) and η Cha (Lawson et al. 2004; Bouwman et al. 2006; Sicilia-Aguilar et al. 2009) show that, despite the differences in environment and disk classes, there is a clear trend for evolved disks to display reduced or even zero accretion rates, consistent with parallel dust and gas evolution. On the other hand, the Hα EW and thus the accretion rates of PTD and full-disks are not significantly different. This is remarkable, since PTD are good candidates to host clean gaps in their disks, maybe related to planet formation (Espaillat et al. 2010). It could be due to difficulties to identify disks with gaps from unresolved observations and lacking spectra of the 10 μm region, or be a direct consequence of most of the full-disks in Tr 37 having moderate accretion rates, compared to younger regions like Taurus (SA06b; Sicilia-Aguilar et al. 2010). Although half of the TD have Hα EW consistent with no accretion, and most of the dust-depleted disks also show relatively low Hα EW, the difference between disked and diskless objects is still remarkable even if TD and dust-depleted disks are included.

We also explored whether the differences in accretion vs. disk type depend on the spectral type. Table 3 reveals that the differences are subtle. Non-accreting TD seem to be more frequent around G/K-type stars. One explanation could be that for a given excess, the holes in a G/K-type star will be much larger than in M-type objects, although it could also indicate that accretion can survive for longer time in the disks around M-type objects, once the inner disk starts dispersing. The data also suggests that M-type dust depleted disks are consistent with no accretion, which is in agreement with our observations of M-type stars in the Coronet cluster (Sicilia-Aguilar et al. 2008, 2011a). Nevertheless, a remarkable difference is that while most of the disks around M-type stars in the Coronet cluster are dust-depleted or have very low disk mass compared to their stellar mass (Md/M∗ < 10^{-4}; Sicilia-Aguilar et al. 2013), strongly dust-depleted disks are a minority among M-type stars in Tr 37, at least considering our detection limits at 24 μm. In general, the differences in accretion and disk structure between G-K type and...

Table 3. Results of the double-sided KS probability test for the Hα EW for objects with different types of disks.

| Disk types       | 2006 sample | 2013 sample | All objects | M-type | G/K-type | Comments          |
|------------------|-------------|-------------|-------------|--------|----------|------------------|
| Full-disks/Diskless | 1E-17       | 6E-8        | 1E-22       | 7E-12  | 2E-11    | Strong difference|
| Full-disks/PTD    | 0.47        | 0.63        | 0.28        | 0.62   | 0.57     | No significant difference |
| Full-disks/TD     | 0.0001      | 0.02        | 5E-6        | 0.009  | 5E-5     | Significant difference |
| Full-disks/Depleted | 0.07       | 0.004       | 3E-4        | 0.001  | 0.24^a   | Significant difference (only M-type stars) |
| Diskless/PTD      | 1E-5        | 4E-4        | 1E-8        | 1E-5   | 5E-4     | Strong difference |
| Diskless/TD       | 0.007       | 0.07        | 0.001       | 0.0001 | 0.28     | Significant difference (only M-type stars) |
| Diskless/Depleted | 0.006       | 0.61        | 0.16        | 0.82   | 0.005^a  | Potential difference (only GK stars) |
| Disked/Diskless   | 1E-14       | 4E-6        | 4E-18       | 2E-10  | 1E-8     | Strong difference |

Notes. The 2006 sample corresponds to the members identified in our previous work, the 2013 sample includes the newly detected objects in this work. ^a Small number statistics, only 6 depleted disks among G/K stars.
4.2. Radiative transfer models of the disks

To extract quantitative information from our previous disk structure classification, we followed the scheme in SA11, using the RADMC radiative transfer code (Dullemond & Dominik 2004) to model a subset of the disks within each class (normal full-disks, PTD, TD, and dust depleted disks). The basic parameters in the RADMC code are the inner disk radius, the outer disk radius ($R_{\text{disk}}$), the disk vertical scale height, and the dust distribution. The inner disk radius is set to the place where the temperature is 1500 K, approximately the dust sublimation radius, and the outer disk radius is taken to be 100 AU, although our data offers a poor constrain on this parameter and there is substantial degeneracy between disk radius and total disk mass. The disk vertical scale height is given as $H/R \propto R^{1/7}$, although we also computed the scale given by hydrostatic equilibrium. Finally, the dust distribution is taken to be a collisional power law with exponent $-3.5$, between a typical minimum size of 0.1 μm and a typical maximum size of 100–10000 μm. The gas to dust ratio is 100, and gas and dust are always assumed to be well-mixed, with all dust grains having the same temperature distribution. The stellar parameters (mass, radius) are derived from the optical data and spectral types, using the transformations in Siess et al. (2000) and Kenyon & Hartmann (1995). The dust is considered as amorphous silicate dust with equal amount of magnesium and iron (Jäger et al. 1994; Dorschner et al. 1995). We also include 25% of carbon grains with the same size distribution than the silicate dust. The SEDs are calculated for an intermediate disk inclination (45 degrees), although for these relatively flat disks, there is little difference unless the disk is viewed edge-on.

As in our previous papers, our models are not aimed at a perfect reconstruction of the SED, but at exploring the parameter space. Starting with the most simplified model possible, we check whether any structural differences are required to obtain SEDs similar to the ones we observe by varying only the disk mass and the vertical scale height (with values not far from hydrostatic equilibrium). Only if no reasonable fit is attained in this way, we proceed to change the inner disk radius, the dust grain distribution (reducing the content of small dust grains or increasing the maximum grain size), and also to introduce radial variations in the vertical structure and/or the dust composition. There is a very strong degeneracy between the flaring of the outer disk, the total disk mass, the grain distributions in the inner and outer disk, and the radius (or disk temperature) at which the disk properties (flaring, density, grain size distribution) change. Constraining the precise properties of individual disks with unresolved observations is highly degenerated, so our purpose is to prove that a simple disk structure cannot reproduce the observed SEDs for TD and PTD, unless some radial changes are included, even if the precise values of the varying parameters cannot be determined. Table 4 summarizes the model disk parameters within each disk class. A brief discussion on each class and the most appropriate models follows.

4.2.1. The typical full-disks

The full-disks (Fig. 9) are easy to reproduce by regulating the disk mass to fit the 24 μm point, and changing the vertical scale
height, either assuming a global power law for the disk flaring at all radii or hydrostatic equilibrium. The required vertical scale heights do not depart much from hydrostatic equilibrium, although we find disks that appear higher than predicted by hydrostatic equilibrium (especially, among the higher mass disks) while other disks tend to be flatter, as expected from dust settling. We modeled three objects, with different spectral types and different disk masses, 213659108+573905636, 213751210+572436151, and 213823950+572736175, in representation of this broad and numerous class.

Disks with large excesses, like 213659108+573905636, require a very large amount of mass in small dust grains. Although the maximum grain size is unconstrained, if we assume a collisional distribution with maximum grain sizes over 100 μm and a gas to dust fraction of 100, we would require disk masses about 10% of the estimated stellar mass. This would be on the upper limit according to submillimeter and millimeter observations (Andrews & Williams 2005, 2007). On the other hand, disks like that around 213823950+572736175, with a low-to-moderate mass, probably require large grains to account for mass enough to sustain the accretion rates we expect from their strong Hα emission.

Other disks have very strong near-IR excesses but a lower 24 μm than expected, suggestive of strong flaring but a very low dust mass, even if we consider substantial grain growth (e.g. 213751210+572436151, 213809997+572352782). This is similar to other disks in the Coronet cluster (Cra-159 and HBC 677; Sicilia-Aguilar et al. 2013). Further observations at longer wavelengths are needed to clarify the total dust content and other special structures that could affect their morphology, like truncation or gaps at larger radii. As a first approach, we consider these disks within the full-disk class.

4.2.2. The pre-transitional disks

The pre-transitional disks (Fig. 10) are the hardest class to identify with our available data. The lack of a mid-IR spectrum and information on the silicate feature are a strong constraint to identify disks with partially optically thin regions in their inner disks (Espaillat et al. 2010; SA11). Nevertheless, the high 8–12 μm fluxes and sharp kinks in the mid-IR suggest disks where the properties in the innermost and outer disk change. Our models do not include gaps, although a gap created by a companion or by photoevaporation can offer a natural barrier to explain the radial changes in properties. Since this class is the most problematic, we selected 6 objects as examples (213658737+573848181, 213734649+571657705, 213945201+574912946, 213954058+572933454) and ran different models to show the degeneration of the various parameters. Our list of models is far from complete, but it shows that, despite the degeneracy, the SEDs reveal that some radial variations in disk properties are needed. The radial changes, reflected in Table 4, are a signature of a distinct inner and outer disk, with different flaring, inner radius, grain sizes, and mass.

The case of 213734649+571657705 shows that a simple full-disk model fails to both reproduce the inner and outer disk properties (dashed line in Fig. 10), while a change in grain size distribution and flaring offers a good compromise to reproduce the SED (bold line). 213945201+574912946 offers a practical example of the impossibility to fit a pre-transitional SED with a simple full-disk model (dashed line), that would require extreme flattening far beyond hydrostatic equilibrium and a very large mass to account for the turn up at 24 μm, but would still overestimate the fluxes at 8–12 μm. A model with a differentiated inner disk and hydrostatic equilibrium comes close to the result...
4.2.4. The dust depleted disks

Dust-depleted disks (Fig. 12) are also relatively simple. They can be reproduced with models without radial variations, but require a very low small dust mass (at least, with very low mass in the submicron to 20–30 μm range) and/or very strong settling. There is a strong degeneracy between mass and vertical scale height. In general, if we assume that the dust-depleted disks are in hydrostatic equilibrium, we require a total mass several orders of magnitude below the typical values of normal full-disks. If we impose strong settling and disk flattening, the small dust grain depletion does not need to be so dramatic, although in general an increase of mass over one order of magnitude (for the same grain size distribution) cannot be compensated by flattening the disk, since the 24 μm fluxes would increase far beyond what is observed (e.g. 213030129+572651433). Some of the depleted disks may also have inner holes; 213854760+572450268 is one example where the failure of various models to reproduce simultaneously the 24 and 8 μm flux may indicate the presence of a small hole. However, this is not always necessary and most of them do not satisfy the criteria for TD.

The mass of the disk also depends strongly on the disk radius and on the maximum grain size, but even assuming millimeter-sized maximum grain sizes requires a very low mass for the disk, compared to typical full-disk values (see for instance 213733557+573550931). This implies that even if our data does not allow us to fully rule out a total higher disk mass in the disks classified as dust-depleted, the mass content in grains smaller than 20–30 μm needs to be very low, which is in any case a strong sign of evolution. From the analysis in Sect. 4.1, the dust-depleted disks clearly show accretion differences with respect to full disks. Therefore, despite the degeneracy between disk mass and disk flattening/settling, which has also been discussed in the literature (e.g. Espaillat et al. 2012), our study suggests that these disks are in a different evolutionary stage from typical full disks. Even if we cannot constrain the mass of the disks, and even if the case is particularly hard for M-type stars that have naturally lower near-IR excesses, the low probabilities that M-type dust depleted disks and M-type full disks are drawn from the same collection of objects, and the fact that M-type dust depleted disks and M-type diskless stars are consistent with the same Hα values (see Table 3) strongly point to a different evolutionary status compared to full disks. Accretion termination/gas depletion may be a sign of parallel dust and gas evolution, and also favor stronger dust settling within the disk. The study of silicate features (as in the case of 01–580; Sicilia-Aguilar et al. 2011a), and far-IR/submillimeter observations are a key to better constrain the disk mass of these objects, their structure, and the presence of settling (see for instance the models for 213930129+572651433 and 213944898+573537212).

4.3. What drives disk evolution and dispersal in Tr 37? Clues from accretion, dust, and environment

Putting together the members found in our previous papers and the newly found ones, we sum up to 361 low-mass (spectral types G, K, M) members in Tr 37. The current results are consistent with our previous estimates of the disk frequency (48 ± 5%, including all objects with excesses down to 24 μm, SA06a), although since this survey was strongly biased towards objects with disks, we cannot re-estimate the disk fraction.
Table 4. Disk models for selected members, according to their disk structure.

| Object      | $T_{\text{eff}}$ (K) | $R_\ast$ ($R_\odot$) | $M_\ast$ ($M_\odot$) | $a_{\text{min}}-a_{\text{max}}$ ($\mu$m) | $H_{\text{radial}}/R_{\text{disk}}$ | $T(R_{\text{disk}})$ (K) | Comments/model reference in plots |
|-------------|----------------------|-----------------------|-----------------------|-----------------------------------------|----------------------------------|------------------------|----------------------------------|
| **Full-disk** |                      |                       |                       |                                         |                                  |                        |                                   |
| 213659108   | 4330                 | 1.49                  | 1.0                   | 0.08                                    | 0.1–100                          | 0.12                    | 1500                             | Very massive, rich in small dust (bold line) |
| 213751210   | 3955                 | 0.80                  | 0.7                   | $2.4E-5$                                | 0.1–10000                        | 0.25                    | 1500                             | Very low-mass, maybe smaller radius? (bold line) |
| 213823950   | 3729                 | 0.92                  | 0.6                   | 2.9E–4                                 | 0.1–100                          | 0.15                    | 1500                             | Normal-to-low disk mass, mass depends on grain size (bold line) |
| "           | 3729                 | 0.92                  | 0.6                   | 2.9E–3                                 | 0.1–10000                        | 0.15                    | 1500                             | Normal-to-low disk mass, mass depends on grain size (dashed line) |
| **PTD**     |                      |                       |                       |                                         |                                  |                        |                                   |
| 213658737   | 3500                 | 1.40                  | 0.3                   | 3.0E–3                                  | 0.1–100                          | 0.13/0.22               | 1500/120                        | Less flared, less dense inner disk (bold line) |
| 213734649   | 3720                 | 1.00                  | 0.45                  | 4.5E–4                                  | 0.1–2/0.1–1000                   | 0.15/0.20               | 1500/300                        | Change in grain distribution inner/out disk (bold line) |
| "           | 3720                 | 1.00                  | 0.45                  | 4.5E–4                                  | 0.1–100                          | 0.16                    | 1500                             | Cannot fit mid-IR and near-IR properly (bold line) |
| 213929250   | 3720                 | 1.00                  | 0.40                  | 2.0E–4                                  | 0.1–2/0.1–1000                   | 0.16/0.18               | 1500/400                        | Radial change in grain size and vertical scale height (bold line) |
| 213929408   | 4060                 | 1.15                  | 0.80                  | 8.0E–4                                  | 0.1–2/0.1–1000                   | 0.13/0.12               | 1500/200                        | Optically thin, small-grain inner disk (bold line) |
| "           | 4060                 | 1.40                  | 0.80                  | 4.8E–4                                  | 0.1–100                          | 0.07                    | 1500                             | An extreme attempt to fit with a uniform disk model (dashed line) |
| 213945201   | 4060                 | 1.40                  | 0.80                  | 4.8E–4                                  | 0.1–100                          | 0.13/0.11               | 1500/200                        | Needs more flaring in inner part (dotted line) |
| 213954058   | 3960                 | 1.10                  | 0.80                  | 4.0E–4                                  | 0.1–100                          | 0.14/0.23               | 1500/150                        | Needs a thin inner disk and a thick, flared outer disk (bold line) |
| **TD**      |                      |                       |                       |                                         |                                  |                        |                                   |
| 213633647   | 5165                 | 2.90                  | 2.00                  | 4.0E–6                                  | 20–1000                          | 0.32/hydro              | 500                              | No small dust, inner hole (dashed line) |
| "           | 5165                 | 2.90                  | 2.00                  | 9.6E–8                                  | 0.1–100                          | 0.32                    | 500                              | Small grains, lower mass, silicate feature (bold line) |
| 213735713   | 3580                 | 1.10                  | 0.35                  | 7.0E–5                                  | 0.1–10000                        | 0.20                    | 250                              | Low mass, low $T(R_{\ast})$, high vertical scale height (bold line) |
| 213756779   | 3650                 | 0.75                  | 0.40                  | 1.6E–5                                  | 0.1–10000                        | 0.15                    | 350                              | Low mass, inner hole (bold line) |
| "           | 3650                 | 0.75                  | 0.40                  | 1.2E–4                                  | 20–10 000                        | 0.08                    | 1500                             | Low scale height, large grains, no hole, too much near-IR excess (dashed line) |
| 213914837   | 3850                 | 1.05                  | 0.50                  | 2.5E–7                                  | 0.1–100                          | 0.06                    | 1500                             | Small grains, very low mass, overpredicts 8–12 $\mu$m excess (bold line) |
| "           | 3850                 | 1.05                  | 0.50                  | 4.0E–6                                  | 20–1000                          | 0.06                    | 1500                             | No small grains, better fit to 8–12 $\mu$m but still high (dashed line) |
| 213954058   | 3960                 | 1.10                  | 0.80                  | 4.0E–4                                  | 20–1000                          | 0.06                    | 800                              | Best fit includes small hole (dotted line) |
| **Depleted**|                      |                       |                       |                                         |                                  |                        |                                   |
| 213733557   | 3700                 | 0.80                  | 0.50                  | 1.5E–6                                  | 0.1–100                          | 0.12                    | 1500                             | Without strong grain growth, very low mass (bold line) |
| "           | 3700                 | 0.80                  | 0.50                  | 1.5E–5                                  | 0.1–10000                        | 0.12                    | 1500                             | Larger grain size, higher mass, but still depleted (dotted line) |
| 213854760   | 3720                 | 1.90                  | 0.40                  | 8.0E–5                                  | 20–1000                          | 0.06                    | 1500                             | Too much 8–12 $\mu$m excess (bold line) |
| "           | 3720                 | 1.90                  | 0.40                  | 2.0E–6                                  | 0.1–100                          | 0.05                    | 1500                             | Extremely low vertical scale height underpredicts 24 $\mu$m (dashed line) |
| 213930129   | 3785                 | 1.69                  | 0.70                  | 7.0E–4                                  | 0.1–100                          | 0.05                    | 1500                             | Hydrostatic equilibrium overpredicts 8–12 $\mu$m excess (dotted line) |
| "           | 3785                 | 1.69                  | 0.70                  | 5.6E–6                                  | 0.1–100                          | 0.08                    | 1500                             | Lower mass disk, close to hydrostatic equilibrium (dashed line) |
| 213944898   | 3500                 | 1.26                  | 0.26                  | 3.9E–6                                  | 0.1–100                          | 0.10                    | 1500                             | Lower mass, higher vertical scale height (bold line) |
| "           | 3500                 | 1.26                  | 0.26                  | 3.9E–5                                  | 0.1–10000                        | 0.10                    | 1500                             | If strong grain growth is assumed, the mass may be higher (black dotted line) |

Notes. The comments also include a reference to the models plotted in the figures. RADMC disk models for selected members with different disk structures (normal full-disk, pre-transitional disks [PTD], transitional disks [TD], and dust-depleted disks [Depleted]). The names of the objects have been shortened to fit the table. The stellar parameters ($T_{\text{eff}}, R_\ast$) are determined from the optical observations and spectroscopy. The disk parameters are modified to reproduce the observed SED in the whole wavelength range, starting with the simplest possible model (see Sect. 4.2). The disk masses are estimated assuming a gas-to-dust ratio of 100 and a collisional distribution for the dust with grain sizes $a_{\text{min}}-a_{\text{max}}$. The outer disk radius is taken to be 100 in all cases. A larger radius would also result in a larger dust mass. For pre-transitional disks with physically different inner and outer disks, we give the grain distributions, vertical scale height, and inner rim temperature for both the inner and the outer disk. In case of different models, the comments include a reference to the appropriate figure.
Among the 218 stars with disks (excluding Class I and uncertain objects), the full-disk objects are the most numerous, with a total of 111 stars (about 51% of the total). The least numerous are the pre-transitional disks (21 disks or nearly 10% of the total) and depleted disks (25 objects, about 11% of the total). The detection rate of PTD is low probably because of the lack of silicate feature observations. Our study of IRS spectra (SA11) suggests that the group of PTD is likely more numerous than estimated here from photometry alone. The group of dust-depleted objects has to be also regarded as a lower limit, since detection of low 24 μm excesses among the latest-type stars are also compromised. Finally, transitional disks sum up to 61 objects, or about 28% of the sample. In total, objects with evidence for inside-out evolution (considering as such both transitional and pre-transitional disks) conform nearly 38% of the total sample, being a substantial fraction of the total population of disked objects. An important degree of dust evolution has thus occurred in Tr 37, also in agreement with the significantly low accretion rates observed among cluster members (SA06b; Sicilia-Aguilar et al. 2010), suggestive of parallel dust and gas evolution in the disks.

Our current spectra do not allow us to estimate the accretion rates of the observed objects due to the uncertainty in the sky subtraction. However, considering as a typical value that of $3 \times 10^{-9} M_\odot/yr$ (Sicilia-Aguilar et al. 2010), one would expect that the disks would have had a minimum disk mass of 0.012 $M_\odot$. To reconcile this value with the typical disk masses observed, in general lower by at least a factor of few to one order of magnitude, strong grain growth needs to be taken into account. This is particularly important in case of the disks with very low dust masses and relatively strong Hα emission (like 213751210+572436151), that would be about to lose their disks within less than 10^4 yr. If we do not include strong grain evolution or an anomalous gas to dust fraction, we would need to assume that are witnessing a very special moment of disk evolution in Tr 37, where a large fraction of disks is expected to...
Fig. 11. Disk models for selected SEDs of transitional disks (magenta lines). See Table 4 for information on the individual models. For comparison, the photosphere of a star with the same spectral type from the MARCS models (Gustafsson et al. 2008) is displayed. All datapoints have been extinction-corrected according to their individual values of $A_V$ and assuming a standard extinction law (see Table A.1). Information about the Hα EW (in Å) is also displayed.

Fig. 12. Disk models for selected SEDs of dust-depleted disks (magenta lines). See Table 4 for information on the individual models. For comparison, the photosphere of a star with the same spectral type from the MARCS models (Gustafsson et al. 2008) is displayed. All datapoints have been extinction-corrected according to their individual values of $A_V$ and assuming a standard extinction law (see Table A.1). Information about the Hα EW (in Å) is also displayed.
disperse within the next $\sim 10^5$ years, which is highly unlikely from the statistical point of view and considering the age spread among Tr 37 members. Dust evolution (and thus a higher disk mass than expected) would be the only way to ensure that formation of giant planets is still possible among Tr 37 disks, if it has not occurred yet.

We also find evidence that not all disks evolve along the same path. Some TD have very large 24 $\mu$m excesses, while others appear to be also dust-depleted, a sign that inner holes can appear both in massive, small-grain-rich, disks and in disks that are substantially depleted of small dust grains. Other disks are depleted of small dust grains but still have near-IR excesses, suggesting that small-dust depletion alone does not trigger the immediate opening of a hole in the disk. Moreover, although dust-depleted disks will also probably suffer inside-out evolution later on and enter the class of dust-depleted TD, they will not go through the “classical” TD phase (with no near-IR excess followed by a sharp increase at mid-IR wavelengths and a strong 24 $\mu$m excess), given that they do not have enough mass nor vertical scale height. This would need to be considered when estimating the lifetime of “classical” TD. The diversity of objects is a sign of the different disk dispersal processes at work in a cluster as old as Tr 37 ($\sim 4$ Myr for the central population). The differences observed between the accretion behavior of full-disks and TD/dust-depleted objects discussed in Sect. 4.1 suggests that gas accretion and dust evolution evolve in a parallel way. However, the fact that about 50% of TD are consistent with no accretion, while this is exceptionally rare among normal full-disks, is also a sign that not all disks with similar SEDs correspond to physically similar objects, confirming our previous results (Sicilia-Aguilar et al. 2010, SA11).

The stars in Tr 37 reveal several paths of disk evolution (see a sketch of disk types in Fig. 13). In general, evolution occurs from the inside-out, but the time scale when an inner hole develops in a disk appears to be different: in some cases, substantial dust depletion, settling, or both, occur without the disk developing an inner hole, while in some other cases, disks develop an inner hole while still being massive and very flared. To open a hole by photoevaporation (e.g. Clarke et al. 2001; Alexander et al. 2006; Gorti et al. 2009), the accretion rate or viscous transport through the disk needs to drop below a certain level first. Therefore, in general objects with no significant accretion, low dust mass, and inner holes are good candidates to photoevaporation-related holes or to very massive companions that can block accretion. Nevertheless, very massive disks would also be less likely to disperse via photoevaporation until their mass and accretion rate have decreased sufficiently to prevent refilling of the inner hole by viscous evolution. This would make it more likely that their holes are related to other processes, like stellar companions or planet formation. Objects with transitional and dust-depleted disks that still have significant accretion are good candidates to suffer strong grain growth (maybe combined with settling), which would affect the IR excess but not the accretion rate, or very low-mass companions, that would not block the flow of gas through the gap. In the case of PTD, the differences in dust properties (especially, grain size) observed between the inner and the outer disk could be due to dust filtering in a planet-formation scenario (Rice et al. 2006), as we had also previously suggested in SA11. Recent spatially resolved observations show that the possibility of dust filtering/dust trapping in certain parts of the disk is a real one, and that gas flow across gaps appears to occur (Van der Marel et al. 2013; Casassus et al. 2013). More data, in particular, silicate feature observations, and ideally also spatially resolved observations, would be needed to better define the class of PTD.

We also explored the dependency of SED shape and disk structure with the spectral type/stellar mass. Figure 14 shows a plot of the median SED for the K-type and M-type objects with disks, scaled in both cases to the flux in band $H$ and corrected by the individual extinction values of the objects. The plot includes the results of all the known members with well-defined IR excesses, excluding those with anomalous SEDs and/or photometric problems. This sums up to 76 K-type stars and 109 M-type stars. Although M-type stars have systematically lower IR excesses over the photosphere and the upper quartile for M-type stars hardly reaches the mean values for the K-type objects, the difference is not as dramatic as observed in the Coronet cluster (Currie & Sicilia-Aguilar 2011). The 24 $\mu$m data in Tr 37 are not complete, so we expect to miss a substantial part of the objects with low 24 $\mu$m excesses, especially among the fainter M-type stars. Nevertheless, the most striking result in the young (1–2 Myr) Coronet cluster is not the presence of disks...
with very low masses, but the absence of massive disks among the M-type stars, which is not the case in Tr 37. The analysis of the Hα EW in the different spectral type classes in Tr 37 also suggest that K- and M-type stars in Tr 37 are more similar in terms of accretion and disk properties than their counterparts in the Coronet cluster. Disks around M-type stars in Tr 37 are apparently less evolved than in the Coronet cluster, although we would expect the opposite from their ages, a signature that time evolution alone cannot explain the observed disk properties in Tr 37 and the Coronet cluster.

Besides age, the main differences between the Coronet cluster and Tr 37 are the number of members and the stellar density. In the Coronet cluster, we find about 50 stars (including resolved multiple systems) within a radius of about 0.15 pc (Sicilia-Aguilar et al. 2013). The membership count is probably close to complete for spectral types between B9 and M6, considering that the cluster has been observed in the optical, X-ray, IR, and submillimeter (e.g. López-Martín et al. 2005, 2010; Groppi et al. 2004, 2007; Peterson et al. 2011; Sicilia-Aguilar et al. 2008, 2011a). Taking into account only the number of stars and not their mass, we arrive to a stellar density of about 3000 stars/pc^3. In Tr 37, if we take a cluster radius of 5 pc, we count about 360 stars with spectral types O6 to M6, including both disked and diskless objects. If we account for a similar number of stars without disks that we may have missed in our surveys, and allow for a binarity rate of 50%, we arrive to about 1000 stars in total, resulting in a density of 2 stars/pc^3. Even if our member number estimates are wrong by a factor of few, the density difference between both regions would still be close to 3 orders of magnitude.

The outskirts of the cluster may be less dense than the cluster center (e.g. Barentsen et al. 2011), but the difference in stellar density between the Coronet and Tr 37 would still be very large. Even if we consider that Tr 37 has suffered substantial expansion to arrive to its current stage (at a typical rate of 1 pc/Myr), an initial cluster size of 2 pc would only produce a density of 30 stars/pc^3. Moreover, Spitzer observations (SA06a) and isochrone considerations (SA05; Getman et al. 2012) suggest that the population in the outskirts is most likely younger and formed in-situ, not by expansion of the original cluster. According to Sect. 7 of Getman et al. (2012) there are 235 (±60) stars down to 0.1 M_☉ near (inside) the IC 1396A globule, within the volume of roughly 75 pc^3 (7 pc^3). Assuming a binary rate of 50%, this gives 5 stars/pc^3 (10 stars/pc^3) near (inside) the globule. This source density is still much lower than that of the Coronet cluster. The immediate result is that, while the early formation phases of stars in a compact, dense cluster like the Coronet would be heavily compromised by star-star interactions that could affect envelopes and disks, the environment in Tr 37 would have been much less interactive, allowing for larger envelopes and also larger disks masses. The initial disk structure may also pre-determine the evolutionary path that the disk will follow. Later on, other phenomena like the effect of the central O6 star may also contribute to disk dispersal in Tr 37 (e.g. Mercer et al. 2009).

### 4.4. The star-cloud interactions in Tr 37: mini-clusters and sequential star formation

Besides the fact that the O6 star HD 206267 shapes the large IC 1396A globule and other structures in the outskirts of Tr 37, our forbidden line images also reveal a smaller-scale interaction between some of the low-mass members and the remaining cloud. Figure 15 reveals a complex shocked structure in the surroundings of V390 Cep (21:36:50.72 +57:31:10.7, also known as MVA-60; Marshall & Van Altena 1987) and the K6 star 14–141 (Sicilia-Aguilar et al. 2004). The star 14–141 is known to show variable accretion and to be associated to variable forbidden line emission suggestive of time-variable shocks in the proximities of the star (Sicilia-Aguilar et al. 2010). In addition, the [S II] images reveal a large area emitting in forbidden lines, associated to the rim of the cloud in the proximity of 14–141. It is not possible to distinguish whether the shock is due to 14–141 or associated to more embedded objects within the cloud. There is a very embedded, reddened object that can be seen in the Spitzer images, although due to the variable background it is not possible to extract accurate photometric information from it.

There is also evidence of a small jet or shock on the other side of the bubble surrounding 14–141, with no object related to it. From its location, it could be related to V390 Cep, which is classified as a variable Ae star (despite its lack of IR excess at any of the Spitzer wavelengths). The shock emission would be located at about 13000 AU projected distance from V390 Cep. In addition, our Hectospec spectra reveal forbidden emission towards a few other stars (213744131+573331130, 213810759+574013683, 213942378+573348653, and 214059633+572210994; see Table A.5). All this suggest that even the low-mass stars in the cluster contribute to shaping the surrounding cloud, which could have an effect on local star formation.

We had already mentioned the age difference between the solar-type stars in the cluster center compared to the most embedded objects to the west of Tr 37, located within the IC 1396A globule (SA05; SA06a; Barentsen et al. 2011; Getman et al. 2012). In this spectroscopic survey, we took advantage of our previous knowledge of the cloud structure to obtain spectra of the objects associated with some smaller globules (of the order of 20′–40′, or approximately 0.1–0.2 pc at a distance of 870 pc; Contreras et al. 2002). The most remarkable structures contain 6 and 2 stars clearly associated with the extended IR emission, respectively, plus other surrounding objects that could be either associated or projected onto the nebula, forming what we call here “mini-clusters” (Fig. 16).

The larger globule or mini-cluster (Fig. 16, left) is remarkable because of the very strong Hα emission of all its associated stars, which have very strong IR excesses as well. This small stellar group had been already noticed in the Hα survey by Barentsen et al. (2011) and in the X-ray/IR study of...
The results of this work are summarized below:...
Acknowledgements. The authors are grateful to V. L. Afanasiev for conducting the SCORPIO observations, and V. V. Kruhshinsky and P. A. Boley for their help in the preparation of the observations with the 6 m Telescope of SAO RAS, to P. Berlind, N. Caldwell, and M. Calkins at the MMT observatory, and to M. Alises, in the preparation of the observations with the 6 m Telescope of SAO RAS, to P. Barentsen, G., Vink, J. S., Drew, J. E., et al. 2011, MNRAS, 415, 103 Bate, M. R. 2012, MNRAS, 419, 3115 Bate, M., & Bonnell, I. A. 2005, MNRAS, 356, 1201 Bessell, M. S., & Brett, J. M. 1988, PASP, 100, 113 Bouwman, J., Lawson, W. A., Dominik, C., et al. 2006, ApJ, 635, 57 Cardelli, J., Clayton, G., & Mathis, J. 1989, ApJ, 345, 245 Casassus, S., van der Plas, G., Sebastian Perez, D., et al. 2010, A&A, 514, L6 Arzoumanian, D., André, P., Didelon, P., et al. 2011, A&A, 529, L6 Espaillat, C., Ingleby, L., Hernández, J., et al. 2012, ApJ, 747, 103 Fedele, D., van den Ancker, M. E., Henning, T., Jayawardhana, R., & Oliveira, J. M. 2010, A&A, 510, A72 Fernie, J. 1983, PASP, 95, 782 Gatti, T., Natta, A., Randich, S., Testi, L., & Sacco, G. 2008, A&A, 481, 423 holes while keeping a relatively massive outer disk, while others lose a significant fraction of their small dust grains before their inner parts are cleared. This points to several effects contributing to disk dispersal.

Our study of Tr 37 does not reveal the striking differences observed the young Coronet cluster, among K-type and M-type stars (Currie & Sicilia-Aguilar 2011). Non-accreting and depleted disks among M-type stars are not as frequent as in the Coronet, even though Tr 37 is substantially older. A possible explanation is the very different environment and, in particular, the stellar density in both regions. While recent studies suggest a density of the order of 3000 stars/pc^3 in the sparse Coronet cluster, the density in Tr 37 is about 2–3 orders of magnitude lower. This suggests that the differences in disk properties observed between both regions could be related to very different environment and initial conditions, like early interactions. Frequent close interactions at a protostellar level could result in lower initial disk masses and a lack of massive disks among low-mass stars, which is what we observe in the Coronet cluster.

Finally, we also studied the stellar content of several small globules in Tr 37, finding that some of them harbor mini-clusters with a few members. The stars within these mini-clusters show preferentially full-disks with strong IR excesses and strong emission lines. This suggests that the objects could be younger than the extended Tr 37 population, which could be indicative of very small-scale sequential (or clumpy) star formation in the region.

References
Afanasiev, V. L., & Moiseev, A. V. 2005, BaZh, 31, 214 Alexander, R. D., & Armitage, P. J. 2009, ApJ, 704, 989 Alexander, R., Clarke, C., & Pringle, J. 2006, MNRAS, 369, 229 Apai, D., Pascucci, I., Bouwman, J., et al. 2003, Science, 310, 834 Appendini, I., Jankovics, I., & Jetter, R. 1986, A&ASS, 64, 65 Arzoumanian, D., André, P., Dufour, P., et al. 2011, A&A, 529, L6 Barentsen, G., Vink, J. S., Drew, J. E., et al. 2011, MNRAS, 415, 103 Bate, M. R. 2012, MNRAS, 419, 3115 Bate, M., & Bonnell, I. A. 2005, MNRAS, 356, 1201 Bessell, M. S., & Brett, J. M. 1988, PASP, 100, 113 Bouwman, J., Lawson, W. A., Dominik, C., et al. 2006, ApJ, 635, 57 Cardelli, J., Clayton, G., & Mathis, J. 1989, ApJ, 345, 245 Casassus, S., van der Plas, G., Sebastian Perez, D., et al. 2013, Nature, 493, 191 Clarke, C., & Pringle, J. 2006, MNRAS, 370, L10 Clarke, C., Gendron, A., & Sotomayor, M. 2001, MNRAS, 328, 485 Cieza, L., Padgett, D. L., Stapelfeldt, K. R., et al. 2007, ApJ, 667, 308 Contreras, M. E., Sicilia-Aguilar, A., Muzerolle, J., et al. 2002, AJ, 124, 1585 Currie, T., & Sicilia-Aguilar, A. 2011, ApJ, 732, 24 Currie, T., Lada, C. J., Plavchan, P., et al. 2009, ApJ, 698, 1 Cutri, R. M., Skrutskie, M. F., van Dyk, S., et al. 2003, YCat.2246, 0C, VizieR On-line Data Catalog: II/246 Dorschner, J., Begemann, B., Henning, T., Jäger, C., & Mutschke, H. 1995, A&A, 300, 503 Dullemond, C., & Dominik, C. 2004, A&A, 417, 159 Dullemond, C., Natta, A., & Testi, L. 2006, ApJ, 645, L69 Espaillat, C., D’Alessio, P., Hernández, J., et al. 2010, ApJ, 717, 441 Espaillat, C., Ingleby, L., Hernández, J., et al. 2012, ApJ, 747, 103 Fang, M., van Boekel, R., Wang, W., et al. 2009, A&A, 504, 461 Fang, M., van Boekel, R., Bouwman, J., et al. 2013a, A&A, 549, A15 Fang, M., Kim, J. S., van Boekel, R., et al. 2013b, ApJS, 207, 5 Fedele, D., van den Ancker, M. E., Henning, T., Jayawardhana, R., & Oliveira, J. M. 2010, A&A, 510, A72 Fernie, J. 1983, PASP, 95, 782 Gatti, T., Natta, A., Randich, S., Testi, L., & Sacco, G. 2008, A&A, 481, 423
### Table A.1. Observations and spectroscopic information.

| Spect. ID | Hα EW (Å) | Li EW (Å) | Sp. Type | $A_V$ (mag) | Disk? | Campaign | Member? | Comments |
|-----------|-----------|-----------|----------|-------------|-------|----------|---------|----------|
| 21339203+573021561 | Abs. | – | X | – | X | 2009-MMT | N | |
| 213541592+573127753 | Abs. | – | X | – | X | 2010-MMT | N | |
| 213542993+573337049 | –16 | – | M4.0 | 1.1 ± 0.6 | Y | 2009-MMT | Y | |
| 213554323+573443434 | –8 | – | Late | – | N | 2009-MMT | N | |
| 213555607+573030406 | Abs. | N | G9.0 | 6.4 ± 1.9 | N | 2010-MMT | N | |
| 213601910+573430689 | –1 | – | K3.0 | 3.1 ± 0.0 | Y | 2009-MMT | X | |
| 213602343+573326502 | Abs. | – | G9.0 | 2.8 ± 0.2 | Y | 2009-MMT | N | |
| 213602762+572834046 | Abs. | – | X | – | N | 2009-MMT | N | |
| 213603888+572712198 | Abs. | – | X | – | pN | 2010-MMT | N | |
| 213606037+573139138 | –0.1 | – | K1.5 | 3.6 ± 0.1 | Y | 2009-MMT | N | |
| 213607980+57367096 | –11: | – | E-Class | – | Y | 2010-MMT | Y | |
| 213610280+573249821 | Abs. | – | K1.5 | 4.1 ± 0.0 | P | 2009-MMT | N | |
| 213611098+572801874 | Abs. | – | K0.0 | 4.6 ± 0.0 | P | 2009-MMT | N | |
| 213615231+572958174 | Abs. | N | K0.0 | 4.6 ± 0.0 | pN | 2010-MMT | N | |
| 213617003+572639925 | –23 | – | M3.5 | 0.8 ± 0.5 | N | 2009-MMT | N | |
| 213617593+572510785 | Abs. | – | X | – | N | 2009-MMT | N | |
| 213617600+573253249 | Abs. | – | K1.5 | 2.7 ± 0.2 | N | 2009-MMT | N | |
| 213618699+573626870 | –4 | 0.3 | M3.5 | 0.9 ± 0.3 | N | 2009-MMT | Y | |
| 213618850+572649552 | –1 | – | K1.0 | 2.8 ± 0.2 | P | 2009-MMT | N | |
| 213619696+573527805 | –6 | – | M4.5 | 1.3 ± 0.8 | N | 2009-MMT | N | |
| 213620628+572839896 | Abs. | X | – | pN | 2010-MMT | N | |
| 213622159+573237898 | –12 | – | M3.5 | 1.1 ± 0.6 | N(24) | 2009-MMT | pN | |
| 213622379+573141225 | –9 | 0.1: | M3.0 | 3.3 ± 0.9 | pN | 2010-MMT | Y | |
| 213624769+572244914 | Abs. | – | K1.0 | 3.4 ± 0.1 | N | 2009-MMT | N | |
| 213625078+572750265 | –80 | – | M4.0 | 1.2 ± 0.6 | Y | 2009-MMT | Y | |
| 213626039+573327463 | Abs. | – | K1.0 | 3.7 ± 0.1 | N | 2009-MMT | N | |
| 213626897+573043451 | –21 | – | M0.0 | 3.6 ± 0.5 | pN | 2010-MMT | Y | |
| 213627742+573022223 | Abs. | N | K3.0 | 5.9 ± 1.8 | pN | 2010-MMT | N | |
| 213636474+573517477 | Abs. | – | K0.5 | 3.0 ± 0.2 | Y(24) | 2009-MMT | Y | |
| 213634910+572612281 | –5 | pN | M1.0 | 1.7 ± 0.6 | pN(8) | 2010-MMT | pN | |
| 213635322+572622430 | Abs. | – | K2.5 | 3.9 ± 1.0 | N | 2009-MMT | N | |
| 213635588+573559377 | Abs. | – | K1.0 | 3.0 ± 0.2 | N | 2009-MMT | N | |
| 213636909+573132683 | –20 | – | K7.0 | 3.3 ± 0.2 | Y | 2009-MMT | Y | |
| 213637760+572458219 | 0 | – | K1.0 | 3.5 ± 0.2 | N | 2009-MMT | N | |
| 213638543+573641056 | –2 | – | M3.0 | 0.4 ± 0.5 | N | 2009-MMT | N | |
| 213639147+572953326 | –3 | 0.2 | G9.0 | 5.3 ± 0.4 | Y | 2009-MMT | Y | |
| 213641152+573702713 | –7 | 0.3 | M1.5 | 2.0 ± 0.7 | N | 2010-MMT | Y | |
| 213641358+572204113 | –11 | 0.1: | M2.0 | 2.0 ± 1.0 | N | 2010-MMT | N | |
| 213642470+572522186 | –27 | – | M3.5 | 1.4 ± 0.7 | Y | 2009-MMT | Y | |
| 213643679+572331935 | Abs. | N | K1.5 | 5.5 ± 1.7 | pN | 2010-MMT | N | |
| 213644390+571387824 | 0 | – | G7.0 | 6.6 ± 0.6 | P | 2009-MMT | N | |
| 213647421+571942074 | Abs. | – | K0.0 | 3.0 ± 0.2 | N | 2009-MMT | N | |
| 213648190+573402084 | –2 | – | M1.5 | 1.4 ± 0.2 | N | 2009-MMT | Y | |

**Notes.** The full table is available at the CDS. This table lists all observed objects, their Li and Hα EW, and membership criteria. The objects IDs are based on the spectrograph (optical) coordinates. If Hα appears in absorption, we label it as “Abs.” in the corresponding column. If Lithium absorption can be excluded for EW larger than 0.1, we label it as “N” in the corresponding column. Membership and presence of disks are marked as “Y” (yes), “N” (no), “P” (probable), and “pN” (probably not). In case the disk excess was observed only at one wavelength, we label it (e.g. P(24) indicates that the object has probably an excess, but only at 24 μm). Undetermined values of the spectral type, uncertain membership, and uncertain IR fluxes are marked by “X”, objects with spectral types earlier than G are marked with “E”. The errors in the spectral type are of 0.5 subtypes for objects K5 or later, 1 subtype for early K stars, and 2–3 subtypes for G-type objects. The best spectral types (in case of multiple values due to different observations with different S/N) are marked with (5). Objects previously identified by Sicilia-Aguilar et al. (2005, 2006a,b) are marked with (3) (and their previous names are given in the comments), those identified by Barentsen et al. (2011) are marked with (2), those in common with Morales-Calderón et al. (2009) are marked with (1), objects with X-ray counterparts in Getman et al. (2012) are marked with (4), and further IR-excess sources in the same paper are marked with (5). In the comments field, we mention if there is any problem due to IR nebular emission (Neb.), nearby objects (Nobj.), if the SED suggest that the object is variable (Var.), if the object appears older than 10 Myr in the V vs. $V – I$ diagram (oVVI), or any other observation regarding the star/disk.
Table A.2. Optical and 2MASS data for confirmed members and probable members.

| Spect. ID | V (mag) | \( R_J \) (mag) | \( I_J \) (mag) | 2MASS ID | \( J \) (mag) | \( H \) (mag) | \( K \) (mag) |
|-----------|---------|----------------|---------------|-----------|---------------|----------|----------|
| 213542993+573337049 | 19.476 ± 0.011 | 17.613 ± 0.030 | 15.506 ± 0.002 | 21354299+573337049 | 14.302 ± 0.037 | 13.439 ± 0.033 | 13.075 ± 0.036 |

Notes. The full table is available at the CDS. *(a)* Anomalous optical colors. *(b)* Near HD 206267. *(c)* Near 21–840, likely contaminated. *(d)* Near bright star, probably contaminated.

Table A.3. *Spitzer* IRAC and MIPS data for confirmed members and probable members.

| Spect. ID | IRAC-1 (\( \mu \)Jy) | IRAC-2 (\( \mu \)Jy) | IRAC-3 (\( \mu \)Jy) | IRAC-4 (\( \mu \)Jy) | MIPS-1 (\( \mu \)Jy) | SED | Comments |
|-----------|-------------------|-------------------|-------------------|-------------------|-------------------|------|----------|
| 213542993+573337049 | 2683 ± 72 | 1982 ± 57 | 1808 ± 222 | 1211 ± 86 | 881 ± 157 | Dep. | |
| 213607980+572637096 | 19.571 ± 94 | 39.189 ± 78 | 103.489 ± 512 | – | 1159.200 ± 6406 | ClassI | |
| 213618699+573626870 | 1251 ± 53 | 835 ± 53 | 708 ± 197 | 146 ± 46 | – | ND | |
| 213619969+573527805 | 522 ± 3 | 343 ± 3 | 257 ± 15 | 163 ± 48 | – | ND | |
| 213622159+573227898 | 2220 ± 53 | 1434 ± 58 | 1244 ± 273 | – | Neb. contamination | |
| 213622373+573141225 | 1669 ± 43 | 1066 ± 55 | 1305 ± 196 | 1984 ± 230 | ND | |
| 213625078+572750265 | 3060 ± 58 | 3013 ± 61 | 3021 ± 224 | 4872 ± 26 | 9196 ± 103 | Full | |
| 213626897+573304351 | 1392 ± 54 | 750 ± 54 | – | – | – | ND | Neb. |
| 213633647+573517477 | 10.624 ± 79 | 6568 ± 51 | 4339 ± 221 | 2982 ± 72 | 1800 ± 310 | TD/Dep. | Var.? |
| 213636099+573132683 | 10712 ± 33 | 10062 ± 77 | 10814 ± 257 | 15797 ± 95 | 13900 ± 3802 | Full | |
| 213639147+572953326 | 156 347 ± 511 | – | 155 812 ± 641 | 150 577 ± 2912 | Full | |
| 213641152+573702713 | 1045 ± 50 | 717 ± 54 | 488 ± 220 | 235 ± 36 | – | ND | |
| 213641358+572204113 | 876 ± 41 | 589 ± 53 | 272 ± 178 | – | – | ND | |
| 213642470+572523186 | 1513 ± 56 | 1220 ± 55 | 1093 ± 201 | 1046 ± 86 | 1525 ± 100 | Full | |
| 213648190+573402084 | 1203 ± 32 | 686 ± 42 | 314 ± 155 | 1809 ± 67 | – | ND | |
| 213653210+572052112 | 732 ± 50 | 395 ± 52 | 673 ± 208 | – | – | ND: 24 \( \mu \)m marginal | |
| 213655201+573030103 | 38.820 ± 158 | 41.537 ± 165 | 41.378 ± 345 | – | 43.405 ± 11034 | ClassI | |
| 213655283+572516668 | 3304 ± 56 | 1803 ± 43 | 1195 ± 252 | 1045 ± 53 | 2166 ± 101 | TD | |
| 213658737+573848181 | 2764 ± 60 | 2294 ± 54 | 2463 ± 205 | 1866 ± 86 | 16507 ± 456 | PTD | |
| 213659108+573905666 | 6278 ± 62 | 6339 ± 53 | 7337 ± 209 | 10397 ± 91 | 47929 ± 108 | Full | |
| 213659472+573134908 | 5918 ± 18 | 5496 ± 50 | 5118 ± 213 | 8387 ± 48 | – | Full | |
| 213701319+573418289 | 1616 ± 53 | 1282 ± 54 | 1091 ± 210 | 887 ± 79 | 6561 ± 106 | TD | |
| 213701401+572442587 | 5219 ± 61 | 3358 ± 61 | 3050 ± 272 | 1278 ± 73 | – | ND: 24 \( \mu \)m marginal | |

Notes. The full table is available at the CDS. Under “SED” we specify the SED type (ND = no disk; ClassI= Class I object; Full= typical full disk; TD = transition disk; PTD = pre-transition disk; Dep. = dust-depleted disk; Unc. = Uncertain SED type). Objects listed as “marginal detection” at 24 \( \mu \)m have probably excesses but their photometry is compromised by low S/N or different sources of contamination.

A3, page 22 of 29
### Table A.4. WISE data for confirmed members and probable members.

| Spect. ID | WISE-1 (mJy) | WISE-2 (mJy) | WISE-3 (mJy) | WISE-4 (mJy) |
|-----------|--------------|--------------|--------------|--------------|
| 213635647+57351747 | 11.54 ± 0.26 | 164.49 ± 3.03 | 169.31 ± 4.54 | 343.75 ± 12.03 |
| 213639147+572953326 | 148.01 ± 3.54 | 164.49 ± 3.03 | 169.31 ± 4.54 | 343.75 ± 12.03 |
| 213642470+572523186 | 1.63 ± 0.04 | 1.27 ± 0.03 | 8.44 ± 0.09 | 2.38 ± 0.99 |
| 213655283+572551668 | 3.68 ± 0.08 | 1.80 ± 0.04 | 5.26 ± 0.10 | 2.56 ±
| 213659108+573905363 | 6.55 ± 0.15 | 6.66 ± 0.12 | 156.60 ± 0.30 | 54.89 ± 1.77 |
| 213701319+573418289 | 1.66 ± 0.06 | 1.25 ± 0.04 | 27.49 ± 0.13 | 10.13 ± 1.65 |
| 213708137+573616213 | 2.36 ± 0.06 | 2.14 ± 0.05 | 28.03 ± 0.19 | 9.78 ± 1.04 |
| 213708879+572107012 | 1.75 ± 0.04 | 1.41 ± 0.04 | 5.01 ± 0.10 | 2.06 |
| 213716349+572640200 | 6.98 ± 0.15 | 7.69 ± 0.15 | 111.07 ± 0.23 | 46.63 ± 1.63 |
| 21372447+5731359 | 2.64 ± 0.07 | 2.48 ± 0.06 | 46.31 ± 0.15 | 7.93 ± 0.93 |
| 213726148+572303562 | 0.39 ± 0.01 | 0.28 ± 0.01 | 10.82 ± 0.10 | 3.12 |
| 213732341+572503204 | 4.03 ± 0.09 | 3.44 ± 0.07 | 50.07 ± 0.17 | 18.94 ± 1.06 |
| 213734649+571657705 | 2.33 ± 0.06 | 2.06 ± 0.05 | 22.18 ± 0.14 | 11.55 ± 0.77 |
| 213738830+572936901 | 13.20 ± 0.29 | 12.68 ± 0.24 | 143.22 ± 0.33 | 30.64 ± 1.21 |
| 213740471+573433203 | 1.66 ± 0.04 | 1.72 ± 0.03 | 72.11 ± 0.23 | 48.74 ± 1.80 |
| 213742167+573431486 | 0.86 ± 0.02 | 0.82 ± 0.02 | 9.69 ± 0.12 | 9.14 ± 0.96 |
| 213747588+573325074 | 59.06 ± 1.14 | 61.98 ± 1.03 | 85.47 ± 1.03 | 155.84 ± 3.59 |
| 213745888+5734521 | 1.00 ± 0.04 | 0.96 ± 0.03 | 3.20 ± 0.13 | 3.93 ± 1.06 |
| 213744543+572200213 | 0.44 ± 0.02 | 0.41 ± 0.02 | 3.40 ± 0.15 | 2.20 ± 0.65 |
| 213746871+573156208 | 1.80 ± 0.05 | 1.66 ± 0.04 | 19.09 ± 0.17 | 11.19 ± 1.01 |
| 213747963+572423232 | 3.98 ± 0.10 | 3.22 ± 0.07 | 56.75 ± 0.24 | 24.11 ± 1.40 |
| 213751210+572436151 | 1.28 ± 0.03 | 1.24 ± 0.03 | 16.08 ± 0.16 | 7.45 ± 0.71 |
| 213756779+573448171 | 0.69 ± 0.02 | 0.51 ± 0.02 | 2.67 ± 0.16 | 3.28 |
| 213809997+573252782 | 2.13 ± 0.05 | 2.13 ± 0.05 | 27.75 ± 0.16 | 2.52 ± 0.77 |
| 213812023+572500774 | 2.09 ± 0.05 | 1.52 ± 0.04 | 19.68 ± 0.13 | 3.97 ± 0.87 |
| 213812641+572033696 | 2.49 ± 0.06 | 2.03 ± 0.05 | 37.05 ± 0.18 | 8.97 ± 0.72 |

**Notes.** The full table is available at the CDS. Only good quality data in agreement with the Spitzer fluxes are taken into account (see text). Uncertain data are marked with **:**. Data that were not used in the SEDs (due to uncertainty or evident contamination) are marked with **a**.

### Table A.5. Other emission lines observed.

| Spect. ID | Line λ (Å) | Obs. I (Å) | EW (Å) |
|-----------|------------|-----------|--------|
| 213659108+57390536 | Hα 6562.0 | 6562.5 | –16.5 |
| " | Hβ 4861.3 | 4860.9 | –22.0 |
| " | Hγ 4010.7 | 4010.6 | –7.2 |
| " | Hδ 3967.7 | 3967.4 | –1.0 |
| " | Ca II 8498.0 | 8498.2 | –4.0 |
| " | Ca II 8542.1 | 8542.1 | –3.7 |
| " | Ca II 8662.1 | 8662.2 | –3.7 |
| 213659472+573134908 | Hα 6562.6 | 6564.5 | –108.3 |
| " | O I 8446.4 | 8446.0 | –4.8 |
| " | Ca II 8498.0 | 8497.9 | –16.8 |
| " | Ca II 8542.1 | 8541.9 | –17.3 |
| " | H I 8598.4 | 8598.2 | –2.0 |
| " | Ca II 8662.1 | 8662.0 | –14.4 |
| " | H I Pa 12 8750.5 | 8750.3 | –1.7 |
| " | H I Pa 11 8682.8 | 8682.8 | –1.7 |
| " | He I 5875.6 | 5875.6 | –2.2 |
| " | He I 6678.2 | 6677.6 | –0.9 |
| " | H I 5874.2 | 5870.4 | –1.9 |
| " | Fe II 5018.4 | 5018.3 | –1.6 |
| " | Fe II 5169.0 | 5168.6 | –2.2 |
| " | Ca II H 3933.7 | 3933.8 | –4.3 |
| " | Ca II K 3968.5 | 3968.0 | –2.3 |
| " | Hα 4861.5 | 4861.7 | –1.8 |

**Notes.** The full table is available at the CDS. Stars with multiple emission lines and their identifications. The lines have been identified according to Sicilia-Aguilar et al. (2012), the NIST database, Appenzeller et al. (1986), and Hamman & Persson (1992). **a** Probably nebular.
Appendix B: SEDs of all the members with IR excesses

The following figures show the SEDs of all the members with IR excesses. Only those with clear excess are displayed; objects with uncertain colors (for instance, due to nebular contamination or contamination by a nearby object) are not displayed, but marked in Table A.3. Objects without an evident IR excess are also not displayed, since they are consistent with bare photospheres. According to the definitions in the main text, the SEDs have been organized by disk types:

- Typical full-disks, with strong IR excesses and evidence of ongoing accretion (see Figs. B.1 and B.2).
- Transition disks, candidates to have inner holes or inside-out evolution, with $[3.6]-[4.5] < 0.2$ and strong excesses at 24 $\mu$m (Fig. B.3).
- Pre-transitional disks, candidates to inner gaps (Espaillat et al. 2010), with reduced near-IR excesses, and strong 8 $\mu$m (and 24 $\mu$m) fluxes (Fig. B.4).
- Dust-depleted disks, candidates to low small-dust mass given their reduced IR excesses at all wavelengths, specially at 24 $\mu$m. Note that some of the transition disks have also very low excesses, so they are also candidates for dust depletion (or candidates for homologously depleted disks; Fig. B.5).
- Class I objects, consistent with very embedded objects with strong accretion and remnant envelopes (all associated to dense parts in the IC 1396A nebula; Fig. B.6).
- Finally, a few objects (usually with very small excesses, uncertain photometry, and/or likely contamination by nearby bright sources; see Fig. B.6) cannot be safely classified in any of the previous groups; they are therefore labeled and excluded from most plots and statistical analysis.

Probable members are labeled as “P”, all other SEDs correspond to confirmed members. In addition, information on the spectral type and the $H\alpha$ emission is also given in the plot, and the data are compared to a MARCS model with the appropriate spectral type (Gustafsson et al. 2008).
Fig. B.1. SEDs of the members and probably members with IR excess typical of full-disks. Inverted triangles represent upper limits.
Fig. B.2. SEDs of the members and probably members with IR excess typical of full-disks (continued). 213917481+571747432 has strong 8 μm, could be PTD. 213911452+572425205 has [3.6]−[4.5] = 0.27 and a strong kink in the SED, it could be a PTD although its 24 μm flux is likely contaminated by cloud emission. 214011348+574414105 has [3.6]−[4.5] = 0.11 but a SED that suggests strong variability and a flared, massive full-disk.
Fig. B.3. SEDs of the members and probably members with IR excess consistent with TD. Objects like 213819411+572203907, 213839571+572916412, 21394010+5730113 may be also dust-depleted. 213750297+570909399 may be partially contaminated by nebular emission but the detection is real. Inverted triangles mark upper limits, and open circles are uncertain values.
**Fig. B.4.** SEDs of the members and probably members with IR excess consistent with PTD. Note that the 24 $\mu$m flux of 213810182+572708407 is contaminated by a nearby star, but the object still presents the kink in the SED typical of PTD.

**Fig. B.5.** SEDs of the members and probably members with low IR excess consistent with dust-depleted disks. This includes objects with disks but no significant 24 $\mu$m detection, in cloud-free regions. Some objects like 213844343+573626211 could be truncated disks. Objects like 213839749+572753080, 213854760+572450268, 213944898+573537212, 214000478+571839617, 214021922+573005424, and 214026262+573402317 may also have inner holes being thus both transitional and dust-depleted disk candidates. Inverted triangles mark upper limits, and open circles are uncertain values.
Fig. B.6. SEDs of the members and probably members with IR excess typical of Class I objects, early-type very embedded stars, and uncertain objects within the cloud. 21374388+5734521 is an uncertain case that could be a depleted disk or a combination of two sources. In the case of 214005669+573202607, and 21448062+572609205, the small excess at 8 μm with uncertain photometry does not allow us to determine if they are TD or diskless sources. These sources are not included in the statistics of disk types discussed in the text.