A new measurement of the cosmic X-ray background

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ABSTRACT

Aims. We present a new measurement of the cosmic X-ray background (CXRB) in the 1.5-7 keV energy band, performed by exploiting the Swift X-ray telescope (XRT) data archive. We also present a CXRB spectral model in a wider energy band (1.5-200 keV), obtained by combining these data with the recently published Swift-BAT measurement.

Methods. From the XRT archive we collect a complete sample of 126 high Galactic latitude gamma-ray burst (GRB) follow-up observations. This provides a total exposure of 7.5 Ms and a sky-coverage of ~7 square degrees which represents a serendipitous survey, well suited for a direct measurement of the CXRB in the 1.5-10 keV interval. Our work is based on a complete characterization of the instrumental background and an accurate measurement of the stray-light contamination and vignetting calibration.

Results. We find that the CXRB spectrum in the 1.5-7 keV energy band can be equally well fitted by a single power-law with photon index $\Gamma=1.47\pm0.07$ or a single power-law with photon index $\Gamma=1.41\pm0.06$ and an exponential roll-off at 41 keV. The measured flux in the 2-10 keV energy band is $2.18\pm0.13 \times 10^{-11}$ erg cm$^{-2}$s$^{-1}$deg$^{-2}$ in the 2-10 keV band. Combining Swift-XRT with Swift-BAT (15-200 keV) we find that, in the 1.5-200 keV band, the CXRB spectrum can be well described by two smoothly-joined power laws with the energy break at $29.0^{+4.1}_{-3.4}$ keV. The CXRB measurements at energies higher than 20 keV, while it is significantly (30%) higher at low energies (2-10 keV).

Conclusions. Taking advantage of both the Swift high energy instruments (XRT and BAT), we produce an analytical description of the CXRB spectrum over a wide (1.5-200 keV) energy band. This model is marginally consistent with the HEAO1 measurement (~10% higher) at energies higher than 20 keV, while it is significantly (30%) higher at low energies (2-10 keV).

Key words. X-Rays:background, X-Rays:diffuse background, Surveys

1. Introduction

The Cosmic X-ray background (CXRB) is usually defined as the integrated emission of all the extragalactic sources in the X-ray energy band (~2-100 keV). The name background comes directly from the first X-ray astronomical observation (Giacconi et al. 1962), when an apparently diffuse background was observed together with the first extra-solar X-ray source (Sco X-1). The CXRB spectral properties, flux and isotropy were accurately measured by combining A2 and A4 experiments on board the High Energy Astronomical Observatory 1 (HEAO1) satellite. The analytical model produced by Gruber et al. (1999), combining A2 and A4 observations with higher energy data has been considered as a reference for many years (G99 model hereafter). Much effort has been spent to quantify the fraction of CXRB emission due to unresolved point sources. As predicted by Giacconi & Zamorani (1987) a combination of large and deep surveys performed by focusing telescopes in the soft part of the X-ray spectrum (< 10) keV has succeeded in resolving almost the entire (80-90%) CXRB, the resolved fraction decreasing at higher energies (Moretti et al. 2003; Worsley et al. 2005). The point sources producing the resolved fraction of CXRB in the 2-10 keV band have been found to be mostly AGN with a small contributions from galaxy clusters and starburst galaxies (Hornschemeier et al. 2006; Bauer et al. 2004; Brandt & Hasinger 2005; Tozzi et al. 2006). Furthermore, a highly anisotropic diffuse component is present at energies lower than 1 keV (Soltan 2007). This is contributed by the Local Bubble, the Galactic halo (Galeazzi et al. 2006) and the intergalactic medium (Cen & Ostriker 1999), while at higher energies, and high Galactic latitude the diffuse component is negligible.

There is a general consensus on the sources from which the CXRB originates, and the background paradox can be considered solved (Setti & Woltjer 1989); nevertheless, the spectrum of the X-ray integrated emission is still very important in the study of the statistical properties of those sources that are too faint to be detected individually by currently operating telescopes, as highly absorbed AGNs and very high red shift quasars. The extrapolation of the AGN observed spectra (unabsorbed and Compton-thin) to the region of the CXRB peak (~30 keV) can
Table 1. A compilation of CXRB flux measurements in the soft energy band sorted by year of publication, compared with the G99 model.

| Instrument       | Flux 2-10 keV [10^{-11} erg cm^{-2} s^{-1} deg^{-2}] | Reference                         |
|------------------|-----------------------------------------------------|-----------------------------------|
| HEAO1            | 1.65±0.17                                          | Gruber et al. (1999)              |
| Rockets          | 2.20±0.20                                          | McCammon et al. (1983)            |
| ASCA-SIS         | 1.92±0.09                                          | Gendreau et al. (1995)            |
| SAX-MECS         | 2.35±0.10                                          | Vecchi et al. (1999)              |
| ASCA-GIS         | 1.94±0.20                                          | Kushino et al. (2002)             |
| RXTE-PCA         | 1.64±0.05*                                         | Revnivtsev et al. (2005)          |
| XMM-Newton       | 2.24±0.16                                          | De Luca & Molendi (2004)          |
| HEAO1-A2         | 1.66±0.08*                                         | Revnivtsev et al. (2005)          |
| Chandra          | 2.19±0.26                                          | Hickox & Markevitch (2006)       |

(*) The original values are 1.91±0.06 and 1.96±0.10 respectively. We correct them, according to the values reported in Table 3 of (Revnivtsev et al. 2005) to account for the cross-calibration with XMM-Newton.

account for ~75% of the peak (Gilli et al. 2007). Compton-thick AGNs are thought to be responsible for the remaining fraction (Treister & Urry 2005; Gilli et al. 2007; Ballantyne & Papovich 2007). Given that even the most recent AGN surveys in the hard band, >10 keV, can add only few percent to this number (Sazonov et al. 2007, 2008; Tueller et al. 2008), the CXRB provides a key boundary condition in the determination of the census of the heavily obscured AGNs. Moreover, an accurate measurement of the CXRB spectrum is an important observational constraint in the study of the very high redshift (z>6) AGNs which will remain unresolved even with the next generation of X-ray telescopes (Salvaterra et al. 2007; Rhook & Haehnelt 2008). Finally, a proper characterization of the CXRB spectrum is also crucial to ensure proper background subtraction in the study of low surface brightness diffuse X-ray emission coming from the outskirts of clusters and groups of galaxies (Gastaldello et al. 2007; Snowden et al. 2008).

Measurements performed in the soft part of the CXRB spectrum with different instruments (see Table 1) yield a scatter which is much larger than the one expected from standard candle flux measurements (Kirsch et al. 2005), meaning that the observed discrepancy cannot be entirely explained by the differences in absolute calibrations of the individual instruments. The large scatter and the poor control on systematic uncertainties in the CXRB measurements led some authors (Ueda et al. 2003; Treister & Urry 2005) to use the G99 model, re-normalized by a factor of ~30%. The underlying (but not verified) assumption is that the G99 spectrum is correct in shape but with the normalization affected by some calibration problems of the HEAO1 instruments. Worsley et al. (2005, 2006) use an even more artificial solution, combining the XMM-Newton CXRB measurement (De Luca & Molendi 2004) up to 8 keV and the re-normalized G99 at higher energies. On the other hand, recently published measurements, performed by means of wide-field not-focused hard X-ray instruments (SAX-PDS, INTEGRAL-IBIS, Swift-BAT) yield results consistent (10% level) with the G99 model in the 20-50 keV range (Churazov et al. 2007; Frontera et al. 2007; Ajello et al. 2008), reversing the recent trend that prefers higher intensities (Ueda 2007).

Here we present a new measurement of the CXRB spectrum in the 1.5-7 keV energy band, obtained by the analysis of the archival data of the X-ray telescope (XRT) on board the Swift satellite (Gehrels et al. 2004), a mission dedicated to the study of gamma-ray bursts (GRBs) and their afterglows. XRT uses a Wolter I mirror set, originally designed for the JET-X telescope Citterio et al. (1994), to focus X-rays (0.2-10 keV) onto a XMM-Newton/EPIC MOS CCD detector (Burrows et al. 2005). GRBs are detected and localized by the Burst Alert Telescope (BAT) Barthelmy et al. (2005), in 15-300 keV energy band and followed-up at X-ray energies (0.3-10 keV) by the X-Ray Telescope. Following-up gamma-ray burst afterglows, the Swift-XRT obtains deep exposures on random positions of the sky, totally uncorrelated with already known bright X-ray sources, providing us with a simple and direct measurement of the CXRB spectrum.

2. Work strategy

For each energy E we can consider the signal S_{tot} registered in a typical high Galactic latitude GRB afterglow follow-up observation, as the sum of 4 factors. These are the GRB signal, the CXRB itself, which is the one we aim to measure, the electronic and particle induced background (NXB or instrumental background) and the stray-light (SL), i.e. the contamination from sources outside the telescope field of view.

S_{tot}(E) = CXRB(E) \times f_{\text{sign}}(E) + NXB(E) + SL(E) + GRB(E). \tag{1}

The GRB afterglow signal can be easily eliminated by filtering data both in space and in time. The CXRB itself is contributed by bright resolved plus faint unresolved sources and it is affected by vignetting (f_{\text{sign}}). We measure the NXB using two different and independent datasets: a two-day observation performed with the focal plane camera assembly (FPCA) sun-shutter closed (SC) and the data collected in a region of the detector which is not exposed to the sky (NES). We evaluate the third element, the SL contamination using a series of off-axis observations of bright sources. Given the high level of the CXRB
isotropy (Shafer & Fabian 1983; Revnivtsev et al. 2008) we can consider this factor as a fraction of the CXRB (SL = \( f_{SL} \times CXRB \)) in such a way that:

\[
CXRB(E) = \frac{S_{tot}(E) - NXB(E)}{f_{vign}(E) + f_{SL}(E)}. \tag{2}
\]

To measure the CXRB spectrum we perform stacked spectral analysis of a large sample made only of GRB follow-up observations. The following four sections are devoted to fully describing the technical details of our work, which allowed us to estimate the terms of Eq. (2) and their uncertainties. In particular, in Section 3 we describe the dataset and the reduction procedures; in Section 4 and 5 we present how we measure the NXB and SL contamination, respectively. In Section 6 we describe how we calculate the vignetting correction and its uncertainty. In Section 7 we present the spectral analysis procedure and its results, discussing them in Section 8.

Throughout this paper, all errors are quoted at 90% confidence level, unless otherwise specified. The photon index is denoted as \( \Gamma \).

### 3. Sample selection and data reduction

During the first 45 months of operation (Jan 2005- July 2008), the Swift-XRT observed some 300 GRB afterglows with typical exposure times of 70-100 ks during the ~ 10 days following the prompt event. We consider all the long (\( T_90 > 2s \)) GRB follow-up observations with a nominal standard exposure longer than 10 ks and Galactic latitude \(|b| > 20^\circ\). Because we find long term variations of the NXB level we consider only data after January 2006 in such a way that the NXB scatter remains lower than 10% (see Section 4). Similar variations (\( \lesssim 3\% \) per year) in the NXB level were observed in SAX-LECS-MECS and ASCA-GIS (Perri & Giommi 2004; Kushino et al. 2002) and were interpreted as due to a gradual drop in the satellite altitude and/or to the cycle of solar activity.

The sample consists of 126 GRB observations from January 2006 to July 2008. For each observation we exclude from our analysis the data collected in the first day (the segment 0) in order to exclude the brightest phases of the afterglows. For each observation we consider only the central 200 pixel radius (7.9 arcminutes) circle, excluding a 30 pixel radius (1.18 arcminutes), corresponding to \( \sim 95\% \) of the encircled energy fraction of a point-like source (Moretti et al. 2005) around the GRB position. This corresponds to a nominal field of view (FOV) of 0.054 square degrees. The real observed sky solid angle varies from observation to observation depending on the precise pointing distribution of the observation.

We reduce data using the standard software (HEADAS software, v6.4, CALDB version Dec07) and following the procedures in the instrument user guide. We replace the standard good time interval (GTI) definition, which is tuned for the observations of bright point-like sources, by more restrictive filters. Due to the failure of the thermo-electric cooler power supply, the XRT CCD temperature is subject both to orbital (\( 4^\circ \) in 5.9 ks) and long term (\( 15^\circ \) on a day time scale) variations, ranging between -70°C and -47°C (Kennea et al. 2005). Dark current and hot pixels are highly temperature dependent and create high instrumental background in the low energy band (0.3-0.7 keV) during observations performed at temperatures higher than \(-52^\circ\)C (Pagani et al. 2007; Moretti et al. 2007). Moreover, due to the low orbit of the Swift satellite, a typical target can be observed no more than 1-2 ks on a single orbit. Therefore the data from single object are split in different segments. Occasionally some reflected light from the Earth limb significantly increases the very low energy (\(< 0.5 \) keV) background at the end or at the beginning of an observation segment. To reduce these effects, we select intervals with CCD temperature \(< 55^\circ\)C and elevation angle (i.e. the altitude of the observation direction on the Earth horizon) > 40°, instead of the standard -47°C and 30°, respectively. Moreover, we consider only data from observation segments longer than 300 seconds and eliminate the first and the last 100 seconds of each orbital segment. After the complete time-filtering procedure, these procedures typically reduce the effective exposure time to 50% of the standard ones. The total nominal exposure time considered is 7.5 Ms with median value of 40 ks for single observations. The final exposure time distribution of the 126 observations, together with the sky-coverage is shown in Fig. 1: the surveyed area is 7 and 1.3 deg² at 10 and 100 ks respectively.

### 4. Instrumental and particle induced signal

To evaluate the instrumental and particle induced background (NXB) we use two different datasets.

First, we use the two day observation performed between 2007-09-04 18:50:00 UT and 2007-09-06 18:42:00 UT when the instrument Sun shutter (0.38 mm thickness stainless steel of grade 302) was closed due to an improper slew which brought the XRT to point \( \sim 15 \) degrees from the Sun. The instrument automatically closed the shutter in front of the camera. For the next two days the usual XRT observations were performed, but with the shutter closed. We apply to these data the same reduction and filtering procedures that we apply to the sky data. The final exposure time for the instrumental background with the shutter closed is shown in Fig. 1: the surveyed area is 7 and 1.3 deg² at 10 and 100 ks respectively.
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Fig. 3. NXB count rate (per pixel) in the 1.5-7 keV energy band as measured by SC and NES datasets. The linear gradient in the CCD vertical direction is evident.

Fig. 4. The instrumental Ni Kα line flux registered in the 56 observations with exposure times longer than 40ks. The continuous line is the average, while the dotted line is the best linear fit. The total scatter of these measurements with respect of the average is at the level of 10%(1σ), as shown in the right panel of the figure.

closed is 67 ks with an average count rate of 1.92±0.05×10⁻⁷ (7.71±0.3×10⁻⁸) counts s⁻¹pixel⁻¹ in the 0.3-10 (1.5-7.0) keV energy band. For the remainder of the paper we will refer to this dataset as shutter closed (SC) dataset.

The second dataset is provided by the data collected in the regions of the detector which are not exposed to the sky (NES). These are four different regions (2507 pixels each) close to the CCD boundary and delimited by the FOV and the corner sources (Fig. 2). The FOV and corner source region definitions are reported in the standard calibration file (CALDB) swxregion20010101v003.fits. In particular the nominal field of view of the telescope is the 300 pixel radius circle centered on the detector pixel (300,300). We conservatively adopted a wider definition of the FOV which is the 323 pixel radius circle centered in the detector pixel (307,300). Then, we define the NES as the parts of the box centered in detector pixel (307,300), width 436 and height 596 lying outside the conservative FOV. The signal registered in these regions has been telemetered since June 2008 when the telemetered detector area was increased to 600×600 pixels. We use all the available data present in the Swift-XRT archive of the photon counting (PC) observations between June-July 2008. This results in a total exposure of 2.4 Ms. The uncertainty in the determination of NXB is one of the main sources of uncertainty of this measurement. For both these datasets we assume that the signal is contributed only by the particle induced and the pure instrumental background. Possible sources of systematic errors are the following. First, the fluorescence from the shutter itself: if the shutter produces some fluorescence lines, our NXB estimate would be systematically higher than the correct value. Second, significant inhomogeneities in the CCD response or in the intrinsic fluorescence background, could bias the NXB measurement in the NES regions. The third error source is the time dependence of the NXB. As already mentioned, in some previous missions, where the NXB background has been estimated by means of observation of the dark Earth, a time dependence on time scale of years has been observed.

To check our data for these systematic errors, we first verify consistency between the two datasets, SC and NES. Comparing the two datasets we find that the NXB in the 1.5-7 keV band displays a gradient in the vertical direction of the CCD, with the bottom regions being 30% lower than the top regions. This trend is very well described by a linear fit with SC and NES datasets being highly consistent (Fig. 3). The consistency of the two datasets give us good confidence that the first two sources of systematics are negligible for our purposes. We do not have a direct way to monitor the NXB time dependence during all the observations (NES data started to be telemetered only in June 2008) over the entire energy band. Nevertheless we can use the data in the 7-8 keV interval, where the CXRB signal is low and the detected signal is dominated by the Ni Kα line at 7.48 keV produced by the fluorescence of the telescope material. This is uniformly distributed over the detector area with a typical count rate of 3.8×10⁻⁸ count s⁻¹pixel⁻¹. We compare the Ni line observed in the NES regions with the one observed during the SC and sky observations. To minimize the statistical error we consider the 56 observations longer than 40ks. We model the data

\[ \text{Ni Kα rate} = \text{3.31+/−0.03 \ times \ 10^{-8}} \]
Fig. 6. Variations of the effective area as function of the off-axis angle, usually called vignetting. The dashed line represents the standard calibration function, while the black points are the values we find for the 2-3 keV energy band.

in the 6.8-8.2 keV energy band by means of a Gaussian plus a straight line. We find that the Ni line flux decreases slightly with mission time (correlation at 2.5σ confidence) producing scatter of ~10% with respect to the average (Fig. 4). We note that the line flux registered in SC and NES data is perfectly consistent with the average of the observed fluxes. For the purpose of the stacked spectral analysis we account for this uncertainty in the final error budget as explained below in Sect. 7.

5. Stray-light

The other main component of the non-cosmic background is the SL. Ray-tracing simulations indicate that this is produced by photons coming from sources that are outside the telescope FOV at distance between 20 and 100 arcminutes from the optical axis of the telescope, whereas the FOV of the telescope (mirror + detector) has a radius of ~12 arcminutes. A fraction of the photons produced by these sources reach the detector after only one reflection on the mirror or even directly, passing through the mirror shells without any interaction. Some X-ray telescopes mount baffles on top of the mirrors to prevent such a contamination. This is not the case for XRT, for which the SL is a significant fraction of the diffuse radiation registered on the CCD.

In order to evaluate the level of contamination in XRT images, we take advantage of the many Crab Nebula calibration observations. Then a series of observations at 7 different off-axis angles, ranging from 15 to 90 arcminutes were performed. We compare them with the on-axis calibration observation. We calculate the fraction of the source flux present on the central 200 pixel radius circle of the detector as the ratio between the flux observed at each distance from the optical axis and the on-axis flux. Given the large off-axis angles the dimensions of the Crab nebula are negligible for our purposes. The results of this analysis are shown in Fig. 5. We split our analysis in several different energy bands, finding no significant variation as function of energy up to 5 keV. These observations clearly show that XRT images are contaminated by the emission of sources outside the field of view up to ~70 arcminutes. Due to the isotropy of the CXRB we can calculate the expected contamination as the surface integral of the relative flux produced by the Crab observed at different off-axis angles. The result of the integration in the 1.5-7 keV band is the CXRB fraction $f_{\text{SL}} = 0.268 \pm 0.015$ (see Eq. 2).

To give an idea of the SL contamination in absolute terms, using the LogN-LogS calculated by Morre...
7. Spectral analysis

The XRT nominal energy band is from 0.3 to 10 keV. The fraction of \( S_\text{tot} \) due to non-cosmic background (NXB+SL) during a typical XRT observation (without any bright source in the FOV) depends on energy and is shown in Fig. 7. In the energy range between 1 and 6 keV the NXB contributes \( \sim 50\% \) of the total signal registered during the observation. As a comparison, in the XMM-Newton data the cosmic fraction of the total signal is 20% (De Luca & Molendi 2004). Above 7 keV the NXB and, in particular, the particle induced background becomes dominant, due to the presence of the Ni (K\(_\alpha\)) and K\(_\beta\) at 7.478 and 8.265 keV) and Au (L\(_\alpha\)) at 9.713 keV) fluorescence lines. We only consider the energy band between 1.5 and 7 keV, excluding the data with energy less than 1.5 keV because the Galaxy contribution is not negligible (Kushino et al. 2002; Hickox & Markevitch 2006).

We account for the vignetting by modifying the nominal ancillary response file (ARFswxpc012s0620010101v010.arf) by the \( f_{\text{vign}}(E) \) factor (see Section 6). We modify the ARF file also to account for the SL contamination according to Eq. (2) with the results reported in Section 5. Finally, we calculate the contamination of the GRB afterglow residuals outside the 30 pixel radius using the analytical PSF model (Moretti et al. 2005). The GRB residual signal is energy dependent and contributes a maximum of 1.1% below 2 keV, becoming negligible above 3 keV. We account for this contamination applying another small energy dependent correction to the ARF file. To summarize, for each energy \( E \) we modify the nominal ARF according:

\[
\text{ARF}'(E) = \text{ARF}(E) \times \left( f_{\text{vign}}(E) + f_{\text{SL}}(E) \right) \times (1 - \text{GRB}_{\text{res}}(E)).
\] (3)

In order to account for the CCD defects and the excluded 30 pixel radius circular region around the GRB position we create an exposure map for each observation. The overall correction, weighted for the exposure time of the single observation, correspond to 6.3% for the sky exposure and 3.2% for the background. We correct this by modifying the BACKSCAL keyword in the spectrum (PHA) files. For the background file we consider the NES regions that provide better statistics than the SC observations (Section 5). The four NES regions are not homogeneous due to the spatial gradient in the NXB. Nevertheless the symmetry of the geometry allows us to use this dataset without any correction. We perform the stacked spectral analysis merging the 126 event files, re-sorting the events and the GTIs, and extracting the spectrum from the 200 pixel radius central circle in detector coordinates. To fit the data and calculate the fluxes we use XSPEC(v12.4).

To account for the systematic uncertainties in the NXB (Section 4), vignetting factor \( f_{\text{vign}} \), Sect. 6) and SL contamination \( f_{\text{SL}} \), Sect. 5) measurement we produce a large number (10,000) of simulated datasets, randomly varying the NXB normalization, the \( f_{\text{vign}} \) and \( f_{\text{SL}} \), according to the appropriate Gaussian distributions. For the NXB normalization we use the mean standard deviation that we observe for the Ni line fluxes in our sample (lower panel of Fig. 5). For the \( f_{\text{vign}} \) and \( f_{\text{SL}} \) we conservatively use a standard deviation equal to 5% which slightly exceeds the estimated errors.

We neglect the Galactic contribution (absorption and emission) and fit our data by means of a simple power law, obtaining the numbers reported in Table 2 and shown in Fig. 8. For the \( f_{\text{vign}} \) and \( f_{\text{SL}} \) we conservatively use a standard deviation equal to 5% which slightly exceeds the estimated errors.

For the f\(_{\text{vign}}\) and f\(_{\text{SL}}\) we conservatively use a standard deviation equal to 5% which slightly exceeds the estimated errors.
3.1. Cosmic variance

To study the variance of our sample we consider the 113 observations with durations longer than 10 ks for which the spectral parameters and the flux of the CXRB can be calculated with an acceptable accuracy (\(\sigma_\chi \sim 15\%\)). We find that the flux distribution is well described by a Gaussian with a standard deviation of \(\sigma_\chi=20.8\%\pm2.4\%\) (upper panel of Fig. 9). The maximum CXRB flux value in our sample, at \(\sim 5\sigma_\chi\) from the mean, is observed in the field of GRB 071028B where the galaxy cluster Abell S1136 is present. This is not surprising, given the fact that Abell clusters are \(\sim 5000\) distributed over \(~27,000\) deg\(^2\) of sky meaning that 1-2 Abell clusters are expected in our survey. Once we account also for the contribution of the statistical error: \(\sigma_\chi^2 = \sigma_\chi^2 + \sigma_\chi^2 = 1.8\%\). We note also that the extrapolation of our simulations to the total surveyed area \((\sim 7\) deg\(^2\)) tells us that the stacked analysis uncertainty due to cosmic variance is negligible.

7.2. Data check

We check our data against any bias due to the time of the observation, CCD temperature, and Galactic latitude. As explained in Section 4 we find a slight dependence of the NXB on the time of the observations. However, as the NXB dependence is small and the NXB contribution is minor the 113 CXRB flux measurements do not have any significant correlation with the observation time (see lower panel of Fig. 10). As already said (Section 3) the XRT CCD temperature is variable due to the fact that it is only passively controlled. Because dark current and hot pixels are temperature dependent, we also plot the 113 flux measurements against the average temperature of the observations (see central panel of Fig. 10). Finally we check the flux measurements against Galactic latitude to exclude any significant contribution from our galaxy to the XRT measurements (upper panel of Fig. 10). For all the three datasets we find that the best linear fit is consistent with a constant at the level of 1 \(\sigma\).

8. Discussion

As mentioned in the Introduction, the CXRB spectrum normalization is still a debated issue.

The Swift-XRT measurement, we present here, is very close to XMM-Newton (Table 1). This is not unexpected as the Swift-XRT effective area calibration has been slightly modified to match XMM-Newton by means of simultaneous observations. For what concerns cross-calibration, Swift-XRT measures fluxes 5-10\% lower than RXTE-PCA during simultaneous observations of 3c273 (Godet et al. 2008). Cross-calibration observations of 1E 0102.2-7219, the brightest supernova remnant in the Small Magellanic Cloud, recently showed that Chandra-ACS, XMM-Newton-MOS, Suzaku-XIS and Swift-XRT agree to within \(\pm 10\%\) for all instruments (Plucinsky et al. 2008). Therefore, the differences with HEAO1 and RXTE-PCA measurements cannot be entirely explained by the absolute calibration differences, as already pointed out by Frontera et al. (2007).

In the region of the CXRB peak \((\sim 30 \text{ keV})\) all the measurements show a acceptable agreement \((\sim 10\%)\), in the soft band the XRT measurement confirms conclusions: in the 2-10 keV band narrow-field focusing telescopes measure CXRB values which are significantly higher than the ones found by wide-field not focusing telescopes. However, the XRT data, although inconsistently higher than the G99 model, smoothly join the higher energy data as we show by the good fit to the XRT and BAT.

Below 60 keV, the G99 model consists in a CPL with \(\Gamma=1.29\) and energy break 41.13 (note that no uncertainties are reported in

\[\text{Fig. 10. CXRB measures plotted against time, Galactic latitude and CCD temperature. Here the plotted error are only the statistical ones. We fit all the three datasets with a straight line (y=qx+a), finding that they are consistent with a constant (q=0) at the level of 1 \(\sigma\).} \]

\[\text{\(E_{\text{peak}} = E_B \cdot \left(\frac{2 - \Gamma}{\Gamma - 1}\right)^{1/\Gamma} = 22.4 \pm 0.4 \text{ keV.} \quad (5)\)}\]
Table 2. The best fit results for the three different models we use. Among brackets are reported the statistical contribution to the total error budget.

| Model | Range [keV] | Norm. [keV cm\(^{-2}\)s\(^{-1}\)sr\(^{-2}\)] | Ph.ind.(Γ) | Flux 2-10 [erg cm\(^{-2}\)s\(^{-1}\)deg\(^{-2}\)] | χ\(^2\)(dof) |
|-------|-------------|---------------------------------------------|-------------|---------------------------------|-------------|
| PL    | 1.5-7       | 3.69\(^{+0.20(0.18)}_{-0.20(0.18)}\)×10\(^{-3}\) | 1.47\(^{+0.07(0.06)}_{-0.07(0.06)}\) | 2.18\(^{+0.12(0.02)}_{-0.13(0.02)}\)×10\(^{-11}\) | 209.1(175) |
| CPL   | 1.5-7       | 3.70\(^{+0.20(0.18)}_{-0.20(0.18)}\)×10\(^{-3}\) | 1.41\(^{+0.06(0.06)}_{-0.06(0.06)}\) | 2.13\(^{+0.13(0.02)}_{-0.13(0.02)}\)×10\(^{-11}\) | 209.7(175) |
| 2SJPL | 1.5-200     | 0.109\(^{+0.03}_{-0.03}\)keV                 | 1.40\(^{+0.02}_{-0.02}\), 2.88\(^{+0.04}_{-0.05}\)×9.0\(^{+0.5}_{-0.5}\) | 2.21\(^{+0.07}_{-0.07}\)×10\(^{-11}\) | 200.5(193) |

Fig. 11. Upper panel: A compilation of flux measurements both in the soft and hard energy bands. For the clearness of the plot, not all the soft energy measurements reported in Table 1 are shown here. Because Gruber et al. (1999) do not report the uncertainties in the best fit values, we use a fiducial 5% error for G99 model. Lower Panel: Ratio of the flux measurements plotted in the upper panel with our joined XRT+BAT fit. Colors are the same of upper panel. For comparison, with the dotted line we plot (only in the bottom panel) the ad-hoc model [Worsley et al. (2005)] used to calculate the resolved fraction.
In summary, breaking the paradigm that G99 spectrum has the correct shape shows that CXRB data can be analytically described by a 2SJPL with the values reported Table 2 and a peak in the energy distribution at 22.9 ± 0.4 keV. In the 1.5-50 keV range, this function is very similar to a CPL with the energy break of 41.13 keV and photon index of 1.4.

We note that the 2-10 keV CXRB flux measurement directly affects the calculation of the CXRB resolved fraction. Moretti et al. (2003), combining shallow and deep surveys and integrating the source number counts, estimate that the resolved fraction of 2-10 keV CXRB is 87 ± 6%. Worsley et al. (2005) refined this calculation finding that the resolved fraction ranges from 80% in the 2-4 keV band to < 60% for energies higher than 6 keV. The main reason for the inconsistency between the two results is the value of the CXRB total flux. Moretti et al. (2003) used an average of a sample of CXRB measurements, yielding a value of 2.02 ± 0.11 × 10⁻¹¹ erg cm⁻² s⁻¹ deg⁻² which is 10% less than the present measurement. As already mentioned, Worsley et al. (2005) used an ad hoc model, combining the XMM-Newton measurement with a re-normalized G99 model. As shown in the Fig.1 (bottom panel) this model, although not motivated from an observational point of view, is not very far from our best fit. If we assume the present measurement for the CXRB and integrate the LogN-LogS of Moretti et al. (2003), we find a result for the CXRB resolved fraction which is 79 ± 6% in the 2-10 keV band, in very good agreement with the average value quoted by Worsley et al. (2005). The values relative to the single narrow bands at higher energies, on the other hand, should be slightly corrected, applying our CXRB value.

The LogN-LogS extrapolation at very low fluxes (10⁻¹⁷ erg cm⁻² s⁻¹ deg⁻², a factor 20 lower than the faintest Chandra deep field sources) cannot account for all the CXRB. This implies that a not negligible fraction of the CXRB is supposed to be produced by non detected sources. Worsley et al. (2009) and Hickox & Markovich (2007) correlate almost the entire CXRB unresolved fraction to optical/IR detected galaxies in the Chandra deep fields. These are star-forming galaxies which are expected to overwhelm the number of AGNs at very low fluxes (Ranalli et al. 2003, Bauer et al. 2004) absorbed AGN (Teister & Urry 2005; Gilli et al. 2007) which are supposed to be the main component at higher energies and with a small contribution from very high redshift (z>6) quasars.

9. Conclusion

We use the Swift-XRT archival dataset to determine the flux and spectrum of the CXRB. This has two main advantages. The first one is the observational strategy which provides us with a truly random sampling of the X-ray sky, not correlated with previously known sources. The second is the low level of the NXB background, which allows measurement of the CXRB with high accuracy. Similar to other focusing telescopes, we find that CXRB flux is significantly higher than HEAO1/G99 model. Nevertheless combining our dataset with Swift-BAT data, we show that we can describe the CXRB spectrum with a simple model (two smoothly joined power laws) over a wide energy band. The model we propose is much more observationally motivated than the ones recently used in the literature for population synthesis models and for the CXRB resolved fraction calculation. Using the present CXRB measurement we calculate that the resolved fraction in the 2-10 keV energy band is 79 ± 6%.

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