On the CCD Calibration of Zwicky galaxy magnitudes & The Properties of Nearby Field Galaxies

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ABSTRACT

We present CCD photometry for galaxies around 204 bright ($m_Z < 15.5$) Zwicky galaxies in the equatorial extension of the APM Galaxy Survey, sampling and area over 400 square degrees, which extends 6 hours in right ascension. We fit a best linear relation between the Zwicky magnitude system, $m_Z$, and the CCD photometry, $B_{CCD}$, by doing a likelihood analysis that corrects for Malmquist bias. This fit yields a mean scale error in Zwicky of 0.38 mag per magnitude: ie $\Delta m_Z \approx (0.62 \pm 0.05) \Delta B_{CCD}$ and a mean zero point of $< B_{CCD} - m_Z > = -0.35 \pm 0.15$ mag. The scatter around this fit is about 0.4 mag. Correcting the Zwicky magnitude system with the best fit model results in a 60% lower normalization and 0.35 mag brighter $M_*$ in the luminosity function. This brings the CfA2 luminosity function closer to the other low redshift estimations (eg Stromlo-APM or LCRS). We find a significant positive angular correlation of magnitudes and position in the sky at scales smaller than about 5 arcmin, which corresponds to a mean separation of $120 h^{-1}$ Kpc. We also present colours, sizes and ellipticities for galaxies in our fields which provides a good local reference for the studies of galaxy evolution.

Key words: Galaxies: Evolution ; Galaxies: Clustering ; Cosmology ; Large-scale structure of the universe

1 INTRODUCTION

Some important aspects of galaxy evolution can only be understood by studying the statistical properties of nearby field galaxies, in particular its luminosity function (LF). As well as providing vital information for galaxy evolution studies, an accurate knowledge of the present day LF is needed to normalize the number counts of galaxies at fainter magnitudes, and to understand the clustering and large scale structure of the galaxy distribution. Moreover, the colour distribution of the nearby galaxies provides a basic reference to determine star formation rates in galaxies.

Much recent work has been directed towards studying the evolution of the LF using samples of faint galaxies at high-redshift, but a consideration of the ensemble of available estimates of the LF at low redshifts suggests that there are still large inconsistencies which must be reconciled before reliable conclusions can be drawn about the implications of deep surveys.

It is common practice to fit the LF to the so called Schechter (1976) form:

$$\phi(L) = \phi_* \left( \frac{L}{L_*} \right)^\alpha \exp \left( - \frac{L}{L_*} \right).$$

were luminosity is related with magnitude in the usual way, $\frac{L}{L_*} = 10^{0.4(M_* - M)}$. Determinations of the low redshift $B$-band LF are available from the Stromlo-APM Survey (SAPM, Loveday et al., 1992), the Southern Sky Redshift Survey (Da Costa et al., 1994), the CfA2 redshift survey (Marzke et al., 1994), and more recently from the ESO Slice Project (Zucca et al., 1997) and the 2dF Galaxy Redshift Survey (Folkes et al., 1999). Determinations in the $R$-band are available from the Las Campanas Redshift Survey (LCRS, Lin et al., 1994) and from the Century Survey (CS, Geller et al., 1997). The best determinations of the Schechter function parameters from each of these surveys are listed in Table 1. The SSRS2 is based on the ESO-Upplands photometry of Lauberts & Valentijn (1989), but transformed to the $B(0)$ system of Huchra (1976). Da Costa et al. (1994) ascribe the difference in the LFs of the two surveys to the large colour-term used by Lauberts & Valentijn to relate the two systems, and so we therefore use $m_Z$ for Zwicky magnitudes and $B(0)$ for SSRS2 magnitudes. Da Costa et al. (1994) quote the $rms$ difference between $B(0)$ and $m_Z$ as
magnitude zero-point would not correct for the apparent discrepancy in the normalisation, which would suggest that a more subtle effect is present in the CfA2 data.

In this paper we present the results of a photometric survey of bright galaxies in the overlap region of CfA2 and the equatorial extension of the APM Galaxy Survey (Maddox et al. 1990a,b; Maddox et al. 1991), which we will use to investigate this apparent discrepancy. In Section 2 we begin by comparing the APM photometry with the Zwicky measurements. We describe our CCD observations in Section 3, and compare our observations with Zwicky’s photometry in Section 4. In Section 5 we give some other properties for the galaxies in our fields and discuss the implications for the LF. In Section 6 we compare our results with the findings of other authors and discuss the possible implications for other results. No direct CCD calibration has yet been published for this part of the APM Galaxy Survey. We will present a detailed comparison of our CCD photometry with the APM data down to the survey completeness limit in a separate paper.

2 A COMPARISON OF ZWICKY AND APM MAGNITUDES

From the information used to calibrate the APM survey we would expect $B \approx b_{APM}^{\ast} + 0.2$, given the mean $(B-V) \sim 0.7$ and the colour equation $b_{APM} = B - 0.28(B-V)$ (Blair & Gilmore 1982; Maddox et al. 1990b; although this shift could increase by as much as 0.07 mag according to the findings of Metcalfe, Fong & Shanks, 1995). Adopting the relation between $m_{2}$ and $B$ used by Efstathiou, Ellis & Peterson (1988), this transformation would imply

$$m_{2} \approx b_{1} + 0.5,$$

which is close to the shift in $M_{\ast}$ adopted by Lin et al. (1996), and, if taken as they stand, would suggest that the results of the two surveys might be consistent, at least in $M_{\ast}$.

We investigated the usefulness of this relation by considering the subset of ~100 galaxies found in the part of the APM Galaxy Survey equatorial extension which overlaps with the Southernmost region of CfA2 (The S+3 sample of Marzke et al., 1994). This sample is drawn entirely from volume $V$ of the Zwicky catalogue. We have used the publicly available CfA1 catalogue as our source for Zwicky galaxies, with magnitudes corrected as described in CfA1. The CfA1 includes redshifts only for galaxies with $m_{2} < 14.5$, but has magnitudes and positions for $m_{2} < 15.5$. Marzke et al find

| Survey | Band | $N_{gal}$ | $\bar{z}$ | $M_{\ast}$ | $\alpha$ | $\phi_{\ast}$ |
|--------|------|----------|----------|-----------|---------|-------------|
| CfA2   | $m_{2}$ | 9063 | 0.025 | $-18.8 \pm 0.3$ | $-1.0 \pm 0.2$ | $0.040 \pm 0.01$ |
| SSRS2  | $B(0)$ | 2919 | 0.025 | $-19.5 \pm 0.08$ | $-1.2 \pm 0.07$ | $0.015 \pm 0.003$ |
| SAPM   | $b_{1}$ | 1658 | 0.050 | $-19.5 \pm 0.13$ | $-1.0 \pm 0.15$ | $0.014 \pm 0.002$ |
| ESP    | $b_{1}$ | 3342 | 0.1  | $-19.6 \pm 0.07$ | $-1.2 \pm 0.06$ | $0.020 \pm 0.004$ |
| 2dF    | $b_{1}$ | 5869 | 0.1  | $-19.7 \pm 0.06$ | $-1.3 \pm 0.05$ | $0.017 \pm 0.002$ |
| LCRS   | $r$   | 18678 | 0.1  | $-20.3 \pm 0.02$ | $-0.7 \pm 0.05$ | $0.019 \pm 0.001$ |
| CS     | $K_{IC}$ | 1702 | 0.06 | $-20.7 \pm 0.2$ | $-1.2 \pm 0.2$ | $0.025 \pm 0.006$ |

Table 1. Luminosity Function parameters derived from currently available surveys.

Figure 1. Luminosity function estimations in the SAPM and the CfA2 as a function of absolute magnitudes $M$ in each magnitude system ($b_{1}$ or $m_{2}$). The continuous lines enclose the 2-sigma region in the SAPM estimation whereas the dashed lines enclose the 2-sigma region in the CfA2 estimate.
much smaller than effect shown in Figure 2. It is possible
nitudes in the APM system are therefore expected to be
found the residual zero point errors on individual plates to
original CCD photometry. Maddox et al. (1990b) also used
regions. The equatorial survey extension has been calibrated
of variations in galaxy surface brightness profiles and the

calibrations are less well determined due to a combination


to be representative of the whole of the Southern part of CfA2 (their Figure 3). We inspected this sample
of galaxies on the film copies of the UKST IIIa-J Sky Sur-
vey plates, and examined the image maps as reconstructed
from the APM survey data. We removed from our sample
all those galaxies which had either been broken up into mul-
tiple image components or which were composed of multiple
components which had been merged into a single object by
the APM software. The distribution of these objects in the
$m_\text{z}, b_\text{j}$ plane is shown in Figure 2. A simple least squares fit
to these data with the slope constrained to be unity gives a
zero-point shift of

\[ m_\text{z} = b_\text{j} - 0.6 \pm 0.5, \]

which is more than a magnitude in the opposite sense to
that implied by equation (2). As will be noted below a detailed


calibration requires a correction for Malmquist bias.

The APM Survey is internally calibrated by matching
images in plate overlaps, with the overall zero-point fixed by
a number of CCD frames distributed over the survey region
(Maddox et al., 1990b). At bright magnitudes ($b_\text{j} \lesssim 17$) the


3 CCD DATA

3.1 Observations

We used the galaxies from the sample discussed in the previ-
ous section as the basis for a CCD survey to investigate the
above discrepancy by providing an independent calibration
for both surveys. This overlap region is essentially defined as

\[ 21^h 50 < \alpha < 3^h 40 \text{ and } -0.25^\circ < \delta < 0.25^\circ. \]

We obtained images with the 2.5m Isaac Newton Telescope (INT) and 1.0m
Jacobus Kapteyn Telescope (JKT) in October 1994. The de-
cision to use two telescopes was motivated by the number
of close groups of Zwicky galaxies in the region, so that we
use the 10$'$ field of view available at the INT to image cluster
fields containing groups of bright galaxies and the 6$'$
field of view of the JKT to image individual galaxies.

We used identical Tektronix 1024 $\times$ 1024 detectors (TEK3 and
TEK4) and Harris $B$ and $R$ filters on the two telescopes.
We obtained observations of 58 fields in two nights at the
INT and 73 fields in three nights at the JKT with exposure
times of 360s and 600s, respectively. We observed a number
of photometric standards from the list of Landolt (1992), at
hourly intervals throughout each night, as well as the field
of the Active Galaxy AKN120 which contains a number of
photometric standards (Hamuy & Maza, 1989).

The data were reduced using standard techniques as im-
plemented in the IRAF ccdred packages, with the exception
that we used a modified version of the overscan correction
routine to compensate for a saturation of the preamplifier
used with TEK4 on the JKT. This effect manifested itself as
a sudden drop in the overscan level following readout of par-
ticularly bright stellar objects in the field, and subsequent
exponential recovery to the normal level as the next $\sim 100$
We used the STARLINK 3.2 Image Detection and Photometry software for our photometric analysis. This is essentially the same as the image detection software used in the construction of the APM Galaxy Survey (Irwin, 1986). The basic input parameters for the detection are the threshold intensity per pixel or surface brightness, $I_0$, and the minimum size of the object to be detected, $A_\text{m}$. Amongst other parameters PISA returns total area $A_T$, the ellipticities and the isophotal magnitude $B_I$ or the corresponding total magnitude $B$ resulting from a curve of growth analysis for each detected object after removal of any overlapping objects.

### 3.2 Image Detection and Photometry

We used the STARLINK PISA image detection and photometry software for our photometric analysis. This is essentially the same as the image detection software used in the construction of the APM Galaxy Survey (Irwin, 1986). The basic input parameters for the detection are the threshold intensity per pixel or surface brightness, $I_0$, and the minimum size of the object to be detected, $A_\text{m}$. Amongst other parameters PISA returns total area $A_T$, the ellipticities and the isophotal magnitude $B_I$ or the corresponding total magnitude $B$ resulting from a curve of growth analysis for each detected object after removal of any overlapping objects.

![Figure 3. Comparison of total magnitudes $B_{\text{INT}}$ for different isophotes: $I = 26.1$ mag arcsec$^{-2}$ to $I = 24.0$ mag arcsec$^{-2}$.](image)

We also obtained $R$-band images of our standards and target fields, with zero points determined to be $R_{0,\text{INT}} = 24.16 \pm 0.02$ and $R_{0,\text{JKT}} = 21.77 \pm 0.02$. We were unable to determine any significant colour term from our standard stars, consistent with previous experience with this combination of CCD and filters which are designed to be very close to the Johnson–Cousins system.

The zero-points for the two combinations of telescope and detector to be

$$B_{0,\text{INT}} = 24.16 \pm 0.02,$$

and

$$B_{0,\text{JKT}} = 21.77 \pm 0.02.$$

We also obtained $R$-band images of our standards and target fields, with zero points determined to be $R_{0,\text{INT}} = 24.33 \pm 0.1$ and $R_{0,\text{JKT}} = 22.11 \pm 0.02$, but we were unable to determine any significant colour term from our standard stars, consistent with previous experience with this combination of CCD and filters which are designed to be very close to the Johnson–Cousins system.

| $I_0$ (mag arcsec$^{-2}$) | $A_\text{m}$ (arcsec$^2$) |
|--------------------------|--------------------------|
| INT1                     | 26.1                     | 96                      |
| INT2                     | 25.7                     | 34                      |
| INT3                     | 25.4                     | 34                      |
| INT4                     | 25.0                     | 17                      |
| INT5                     | 24.6                     | 10                      |
| INT6                     | 24.4                     | 10                      |
| JKT1                     | 25.0                     | 12                      |
| JKT2                     | 24.5                     | 12                      |
| JKT3                     | 24.1                     | 6                       |
| JKT4                     | 23.7                     | 6                       |
| JKT5                     | 22.9                     | 3                       |

Table 2. Input parameter sets used for image analysis.

Given that our CCD survey was designed to provide a calibration of both the APM and Zwicky data, we were necessarily interested in a wide range of magnitudes ($14 \leq B \leq 20$) and image sizes. For this range of objects there was no unique combination of $I_0$ and $A_\text{m}$ that could deblend the faint objects without breaking the bright ones. In order to automate this process as much as possible for the whole range of magnitudes, we ran PISA several times with different combinations of $I_0/A_\text{m}$. These are listed in Table 2.

We chose a larger (smaller) area for the fainter (brighter) isophotes so as to select similar objects in all runs. Objects in the final catalogue were selected from the INT4 and JKT4 runs. The information from different isophotes was then used to perform an automatic rejection of broken or contaminated images. After rejection, the total magnitude and size were the largest remaining estimates of $B$ and $A_T$, which typically correspond to the faintest isophote left. The error in $B$ was taken to be the variance in the different estimates for total magnitudes. Objects with rejected isophotes were automatically labeled and checked visually. We refer to total magnitudes estimated in this way as $B_{\text{INT}}$ and $B_{\text{JKT}}$ for the INT and JKT sets, respectively, or $B_{\text{CCD}}$ in general. These effectively correspond to total magnitudes detected at isophotes 26.1 mag arcsec$^{-2}$ and 25.0 mag arcsec$^{-2}$, respectively, unless stated otherwise.

To check if our isophotes are low enough we have also computed total magnitudes determined by PISA using higher isophotes. We find a small zero-point shift of the total $B_{\text{INT}}$ scale as a function of the isophote $I$, but for a given $I$ this shift is not a function of $B$ over the wide magnitude range considered here. This is illustrated in Figure 4, which shows a change of $\approx 0.2$ mag in the mean $B_{\text{INT}}(I)$ when we change the isophote from $I = 26.1$ mag arcsec$^{-2}$ to $I = 24.0$ mag arcsec$^{-2}$. Changing the isophote from $I = 26.1$ mag arcsec$^{-2}$ to $I = 25$ mag arcsec$^{-2}$ introduces a change of only 0.1 mag. Changing the detection isophote from $I = 26.1$ mag arcsec$^{-2}$ to $I = 24.4$ mag arcsec$^{-2}$ also introduces a change in $B_{\text{INT}}$ of about 0.2 mag. We therefore conclude that the residuals associated with our final choice of detection isophote are small compared to the shifts illustrated in Figure 4.

This analysis also implies that there should be a mean zero-point shift in the JKT data relative to the INT data of 0.1 mag due to difference in the isophotes used. We therefore
apply this shift to all objects observed with the JKT to transform these data to our $B_{26.1}$ system.

Finally, we investigated the possibility that there could be a small effect due to the change in mean redshift of the samples at fainter magnitudes by searching for a change in the mean galaxy colour. Figure 4 shows $(B_{INT} - R_{INT})$ as a function of $B_{INT}$. There is only a small colour evolution within the errors. A linear fit to the binned data in Figure 4 yields $B - R = 0.016 B + 1.23$ with a mean $B - R = 1.50$ (the mean weighted by each galaxies is $B - R = 1.61$ as is dominated by the faint objects).

4 COMPARISON WITH THE ZWICKY CATALOGUE

In Figure 5 we show the number counts histograms for the 204 Zwicky galaxies in our sample to the different magnitude systems: Zwicky (top), APM (middle), and CCD (bottom). Comparison of the APM and CCD data shown here suggests that the effect shown in Figure 2 is unlikely to be an artefact of saturation effects in the equatorial APM Survey data at bright magnitudes (see Section 3).

We searched for possible correlations between the magnitude differences, $B_{CCD} - m_Z$, and other properties of our sample. The top panel of Figure 6 shows the $B_{CCD} - m_Z$ difference as a function of CCD colour $B - R$. The colour of Zwicky galaxies is similar to the mean colour in Figure 4, $B - R \approx 1.5$ and there is no apparent correlation with $B_{CCD} - m_Z$ within the scatter. This is a useful check, as any large differences that were due to processing errors might be expected to show up as objects of extreme colour. The bottom panel of Figure 6 shows the scatter in $B_{CCD} - m_Z$ as a function of right ascension. Objects were observed in order of increasing RA on each night of our observing run, and so we would expect temporal drifts to show up as a correlation with RA. Again, there is no evidence of any trend. We also investigated possible variations with Galactic latitude (not shown here), and again found no evidence for trends in our data.

The top panel of Figure 7 shows the $B_{CCD} - m_Z$ error as a function of the mean surface brightness for all Zwicky galaxies. Surface brightness is defined here as the ratio of isophotal magnitudes to the area above a threshold of
25.0 mag arcsec$^{-2}$ in the INT (triangles) and with a threshold of 24.1 mag arcsec$^{-2}$ in the JKT (circles). The difference in the threshold explains the systematic shift in the bulk value of mean high surface brightness of the JKT images (as we are closer to the galaxy core). As can be seen in the Figure, there is no apparent variation of magnitude error with surface brightness.

The bottom panel of Figure 7 shows the $B - m_Z$ error as a function of the sequential run number. This is close to observational time for the INT (triangles) or JKT (circles) when considered separately (in practice, some of the INT and JKT observations were made on the same night). There is no apparent variation of magnitude error with run number.

Figure 8 is the main result of this paper. It shows the $B_{\text{CCD}} - m_Z$ as a function of $B_{\text{CCD}}$. The dotted line shows the $m_Z = 15.5$ magnitude limit. The dashed line corresponds to a direct least square fit to the data. The continuous line shows the fit corrected for Malmquist bias (see text).

4.1 Malmquist bias correction

Different magnitude system are subject to different systematic errors, and even in the best of the situations intrinsic differences in the morphology, environment and spectrum of the galaxies tend to introduce stochastic fluctuations in any magnitude system. We want to find a best fit linear relation between the Zwicky, $m_Z$, and CCD, $B_{\text{CCD}}$, system:

$$m_Z = \lambda B_{\text{CCD}} + Z,$$

where $\lambda$ will account for any scale difference and $Z$ is a zero point shift. In general, both $m_Z$ and $B_{\text{CCD}}$ are stochastic variables and equation (6) is just a mean relation. As is common practice we will assume that there is Gaussian scatter around the above mean relation (due to the accumulation of multiple uncorrelated factors). That is, given a measured magnitude $m_Z$, the error is given by:

$$P(\Delta) = \frac{1}{N} \exp \left[ -\frac{\Delta^2}{2\sigma^2} \right],$$

where $\Delta \equiv m_Z - \bar{m}_Z$ is a stochastic variate, $\bar{m}_Z$ is some mean best fit value in the linear relation of equation (6), $\sigma$ is the rms error and $N$ is a normalization factor. For a sample that is not magnitude limited we have $N = \sqrt{2\pi}\sigma$. For a magnitude limited sample, where $m_Z < m_{\text{lim}}^{m_Z}$, the probability is the same but has a different normalization:

$$N = \int_{m_{\text{lim}}^{m_Z}}^{m_Z} dm_Z \exp \left[ -\frac{\Delta^2}{2\sigma^2} \right],$$

because not all magnitudes are possible, so that $\Delta \equiv m_Z - \bar{m}_Z < m_{\text{lim}}^{m_Z} - \bar{m}_Z$. Thus, we can write $P(m_Z)$ in terms of the complementary error function.

Figure 7. $B_{\text{CCD}} - m_Z$ as a function of surface brightness (top panel) and run number or time of observation (bottom).

Figure 8. The magnitude error $B_{\text{CCD}} - m_Z$ as a function of $B_{\text{CCD}}$. The dotted line shows the $m_Z = 15.5$ magnitude limit. The dashed line corresponds to a direct least square fit to the data. The continuous line shows the fit corrected for Malmquist bias (see text).
The scaling relation between the two magnitude systems is
change of magnitude scale. The zero point different
the above analysis indicates both a zero-point shift and a
value,
the error probability of equation 10, because this depends
comparison with the Malmquist bias corrected version for
because of the Malmquist bias, which produces a deficit of
mentioned above the Gaussian does not provide a good fit
result.

We are now able to perform a likelihood analysis to
find the best fit values of $\lambda$ and $Z$ in the linear relation of
equation 10. We define a likelihood as:

$$L = \prod_i P(\Delta_i)$$

where $i$ runs over all galaxies in the survey, and:

$$\Delta_i = m^Z_{||} - (\lambda B_{CCD} + Z)$$

with $m^Z_{||}$ and $B_{CCD}$ the measured Zwicky and CCD magnitudes for galaxy $i$. In analogy with the standard $\chi^2$ test we
define a Malmquist bias “corrected chi-square”:

$$\chi^2_{Malm} \equiv \sum_i \left[\frac{\Delta^2_i}{2\sigma^2} + 2\log(\text{erfc} \frac{\lambda B_{CCD} + Z - m^Z_{||}}{\sqrt{2}\sigma})\right]$$

Note that the measured magnitude errors, $\sigma_i$, are added
in quadrature to the stochastic error in the linear relation, $\sigma$, which is one of the parameters we want to determine with
this fit. The best fit values that we find on maximising the
likelihood, using the data in Figure 8 are:

$$\lambda = 0.62 \pm 0.05$$

$$Z = 5.9 \pm 0.7$$

$$\sigma = 0.40 \pm 0.05$$,

the errors correspond to approximate 99% confidence levels.
The values in each uncertainty range are strongly correlated:
Lower values of $\lambda$ (eg larger scale errors) are (linearly) cor-
related with larger values of $Z$, and both are obtained for
the smaller $\sigma$. The inverse relation gives:

$$B_{CCD} \simeq 1.61 m^Z_{||} - 9.5$$.

Figure 8 shows a continuous line the best fit model for
the corresponding best fit magnitude difference:

$$B_{CCD} - m^Z_{||} = (0.38 \pm 0.02)(B_{CCD} - 15.53) \pm 0.4$$

Figure 8 shows the residuals of this fit. For comparison we
display a Gaussian with same dispersion, $\sigma = 0.4$. As
mentioned above the Gaussian does not provide a good fit
because of the Malmquist bias, which produces a deficit of
faint objects. It is not possible to show in this Figure a
comparison with the Malmquist bias corrected version for
the error probability of equation 18, because this depends not
only on the differences, $\Delta$, but also on the measured
value, $B_{CCD}$.

Thus, after correcting for the effects of Malmquist bias,
the above analysis indicates both a zero-point shift and a
change of magnitude scale. The zero point different $Z_0$ could
be obtained by taking the mean of the magnitude differences
over $m^Z_{||}$ magnitudes (which are the ones that define the
survey limit):

$$Z_0 \equiv B_{CCD} - m^Z_{||} = -0.35 \pm 0.15$$.

which is in good agreement with Efstathiou et al. (1988).
The scaling relation between the two magnitude systems is then:

$$P(\Delta) = \frac{2 \exp \left[ -\frac{\Delta^2}{2\sigma^2} \right]}{\sqrt{2\pi} \sigma \text{erfc} \left( \frac{m^Z_{||} - m^Z_{||}}{\sqrt{2}\sigma} \right)}$$

so that for $m^Z_{||} = \infty$ we recover the standard Gaussian
result.

The distribution of magnitude errors from the mean
linear fit compared to a Gaussian with dispersion $\sigma = 0.4$. The
Gaussian does not fit the distribution because of the Malmquist
bias, which produces a deficit of faint objects.

$$\lambda \equiv \frac{\Delta m^Z_{||}}{\Delta B_{CCD}} \simeq 0.62 \pm 0.05$$.

We can also test the above model for the scale error
by estimating the magnitude correlations as a function of
projected separation. Figure 9 shows the mean correlation $\langle \Delta(\theta_i) \Delta(\theta_j) \rangle$ as a function of the pair separation $|\theta_j - \theta_i|$, in arcmin. The uppermost panel shows (as con-
 tinuous lines) the autocorrelation for magnitude differences
in the Zwicky system, $\Delta m^Z_{||} = m^Z_{||} - m^Z_{||}$, where $m^Z_{||}$ is the
Zwicky magnitude for the $i$th galaxy and $m^Z_{||}$ is the
mean Zwicky magnitude in the Survey. The middle
panel shows the corresponding autocorrelation for the CCD
system: $\Delta m^Z_{||} = B_{CCD} - B_{CCD}$. The lower panel shows the
autocorrelation of the magnitude errors: $\epsilon_i = B_{CCD} - m^Z_{||}$.
The long-dashed in each case show the zero values for ref-
ence. The short-dashed line in the lower panel shows the
prediction based on applying equation 18 to the data, which
is extremely close to the observed result.

We can see in Figure 10 that there is a significant an-
gular correlation between nearby magnitudes and positions
with $\theta \lesssim 5'$. This correlation is followed by the errors, indi-
cating that the above model is valid. The Zwicky system
shows smaller correlations and smaller variance than the
CCD system. This can also be understood in the context
of the model, as in the Zwicky system the “effective” magni-
itude scale is about a factor of two smaller (equation 18).

Note that at the typical depth of the survey, $D \simeq 80 h^{-1}$ Mpc, the above magnitude correlations are only sig-
nificant at physical scales smaller than $\lesssim 100 h^{-1}$ Kpc. This
correlation has little effect on typical galaxy clustering
scales, $r_0 \simeq 5 h^{-1}$ Mpc, but might be relevant for the
inversion of angular clustering on smaller scales. For ex-
ample, Szapudi & Gaztaña (1998) find that on scales $\theta \lesssim 0.1^\circ$ there is a significant disagreement between the
APM and the EDSGC that is attributed to differences in
the construction of the surveys, most likely the dissimilar
debending of crowded fields. At the depth of these Surveys,
$D \simeq 400 h^{-1}$ Mpc, the above magnitude correlations corre-
Figure 10. Mean correlations in magnitude differences $<\Delta(\theta_i)\Delta(\theta_j)>$ for pairs of galaxies separated by distance $|\theta_i - \theta_j|$, in arcmin shown for Zwicky magnitudes (top panel), CCD magnitudes (middle panel), and for the difference $\epsilon$ between Zwicky and CCD magnitudes (lower panel). The short-dashed line shows the prediction using the scale error model $\epsilon = 0.6\Delta B$. The long-dashed line shows zero correlation for reference.

An illustration

An illustration of the scale error in the Zwicky system can be seen in the galaxies shown in Figures 11 and 12. Table 3 gives the Zwicky and CCD magnitudes for each of the Zwicky galaxies as labelled in the Figures. As can be seen from the table, the range of Zwicky galaxies in Figure 11 is $\Delta m_Z = 0.3$ which is almost 6 times smaller than the CCD range $\Delta B_{CCD} = 1.7$. The range for the whole cluster is $\Delta m_Z = 1.3$, almost a factor of 3 smaller than the CCD range: $\Delta B_{CCD} = 3.5$. These illustrations are strongly suggestive of an observer-bias effect, whereby some fainter galaxies are included in the same due to their proximity to brighter objects, e.g. object #9 in Figure 11.

5 THE PROPERTIES OF NEARBY FIELD GALAXIES

We will now present some further properties of the galaxies in our sample: colours, sizes and ellipticities, and discuss the implications for the Luminosity function. This will help us to understand the systematic effects present in the Zwicky magnitude system. As mentioned in §1, these local properties are interesting in the context of galaxy evolution and star formation rates.

There are still only few samples that are both homogeneous and large enough to provide estimates of the statistical properties of nearby samples, most of which have been selected from photographic plates. Besides the redshift catalogues listed in Table 1, one of the more extensive catalogues of bright galaxies is contained in the Second Reference Catalogue of Bright Galaxies (de Vaucouleurs, de Vaucouleurs & Corwin 1976, known as RC2), which gives magnitudes, colours, ellipticity and morphology for well over a thousand galaxies to a limiting isophote of around 25.0 mag arcsec$^{-2}$. The problem with this catalogue is that it is a mere compilation of data and was not intended to be complete to any...
specified limiting magnitude, diameter, or redshift. Moreover it is based on photographic plates.

Our sample is homogeneous enough and extends over a large enough area (≈ 400 square degrees) to provide a fair sample. Thus the new results based on the CCD magnitudes should be a good local reference for the fainter studies of galaxy evolution.

We will compare the properties of galaxies in the Zwicky ($m_Z < 15.5$) sample with the corresponding properties of galaxies in the INT fields selected using the CCD blue magnitude $B = B_{CCD}$. We choose two magnitude cuts: $B < 16$ which includes most Zwicky galaxies and $19 < B < 20$, which includes faint galaxies in the same fields.

5.1 Sizes and Ellipticities

Figure 13 shows an histogram of the frequency distribution of galaxies as a function of its area (number of galaxy pixels above detection threshold in the CCD image). Top panel shows Zwicky selected galaxies only. Middle and bottom panels include all the galaxies in the INT fields selected with CCD magnitudes of $B < 16$ and $19 < B < 20$, respectively.

As can be seen in the Figure, Zwicky selected galaxies have a long tail of small objects which is not present in the bright subsample of $B < 16$ CCD selected objects. This again illustrates the scale error (and selection biases) in the Zwicky system mentioned above (see also § 6).

The distribution of sizes in logarithmic scale for both CCD sub-samples are well approximated by Gaussians (shown as continuous lines in the figure). The mean size is about 60 arcsec$^2$ for $19 < B < 20$ and 1600 arcsec$^2$ for $B > 16$, with a rms dispersion of about 20% and 27% respectively (recall that these areas correspond to a nominal threshold of 26.1 mag arcsec$^{-2}$, fainter thresholds could be needed to sample the size of the lower surface halos).

Figure 14 shows the corresponding frequency distribution for the ellipticities as measure in the galaxy shapes (with a threshold of 26.1 mag arcsec$^{-2}$). These Figures are in rough agreement with the results by Binney & Vaucouleurs (1982) over the Second Reference Catalogue (RC2).

The local distribution seems remarkably similar to the faint one, given the large differences in sizes shown in Fig. 13. This is a good indication that our isophote (26.1 mag arcsec$^{-2}$) is low enough, as otherwise we would expect a large excess of round faint objects. There seems to be nevertheless a slight (≈ 5%) excess of round faint objects. This is probably not due to the seeing (or pixel resolution) as most of the faint objects here have more than 30 arcsec$^2$ (or 100 pixels) of area (see Fig 13). The lack of correlation between ellipticities $\epsilon$ and areas $A$ is illustrated in the bottom panel of Figure 13. A least-square-fit to the points give $\epsilon \approx 0.277 + 0.018 \log_{10}(A)$.

5.2 Colours

Both the morphology and a more detailed study of colours will be presented elsewhere. Here we just show the colour frequencies and discuss some of its implications.

Figure 15 shows the frequency distribution of CCD colour ($B - R$) for all Zwicky galaxies (top) in our sample. We then show the corresponding distributions for the same Zwicky galaxies separated into two sets, corresponding to the JKT (middle) and INT (bottom) CCD frames. As mentioned above, the INT field of view (10') was larger than the JKT (6') and so we used the INT to target groups and clusters of galaxies (which could take up several overlapping CCD images).
Harris filters which were designed to be very close to the colours presented by Fukugita et al. (1995). We have used different morphological types. Previous studies were limited to exclude all the galaxies that are in the same CCD frame, while the JKT was used for more isolated galaxies. As can be seen in the Figure the INT galaxies are significantly redder, \( B - R \approx 1.57 \), than the JKT, \( B - R \approx 1.36 \). There are about \( \approx 100 \) objects in each subset which roughly corresponds to a random sampling of the total 600 Zwicky galaxies in our survey area (the 600 targets were split in half between the two fields and then more or less randomly selected during each night). The mean colour of our 200 Zwicky galaxies is \( B - R \approx 1.47 \). Neither the mean colour or the frequency distribution change significantly when we exclude all the galaxies that are in the same CCD frame (eg to avoid nearby ellipticals). This indicate that selection effects are not important for this distribution.

Given the large area covered (several hundreds of square degree), these results should be a good estimate of the overall mean colour of bright local galaxies. How do these values compare with previous studies? This is a difficult question because it requires both accurate colours and the accurate fraction of galaxies in different environments (eg different morphological types). Previous studies were limited to inhomogeneous samples (eg RC2) or to small numbers of galaxies. For example Kennicutt (1992) presented results for 8 early-type galaxies and 17 spirals to irregulars. Our mean \( B - R \approx 1.47 \) can be compared against the synthetic \( B - R \) colours presented by Fukugita et al. (1995). We have used Harris filters which were designed to be very close to the Johnson–Cousins (with J Johnson) system. In these broadband, Table 3 of Fukugita et al. show \( B/R = 1.67 \) for Ellipticals, 1.48 for S0, 1.2 – 1.1 for Spirals and 0.6 for Irregulars. The range seems in rough agreement with Figure 3, although we find a significant number of objects with \( B - R > 2 \). Notice that these objects are mostly in the INT sample, that is in clusters or groups. While most of the bluer objects with \( B - R < 1 \) are in the JKT fields, that is, they are more isolated (by angular separations larger than 6', which corresponds to a mean separation larger than 140 h\(^{-1}\) Kpc. Nevertheless, this is seems to be a small effect.

Note that Fukugita et al synthetic colours seem to be slightly less red than the Kennicutt (1992) observations (according to Table 2 in Fukugita et al., they are about 0.1 redder in B-V and this could be larger in B-R). Also notice that we are using total magnitudes, while the synthetic colours are based in small aperture spectra. Galaxies could have quite a different colour distribution (or spectra) in their low-surface halos.

Figure 14 shows the frequency distribution of CCD colour \( B - R \) for the Zwicky galaxies (top panel) in comparison with the same galaxies selected with CCD magnitudes: with \( B < 16 \) (middle panel) and a sub-set of INT galaxies with \( 19 < B < 20 \) (bottom panel), which corresponds to the points in Figure 3. The continuous line in all cases shows for reference a Gaussian distribution with mean \( B - R = 1.45 \) and \( \text{rms} \) deviation of \( \sigma = 0.4 \), which roughly matches the Zwicky frequencies.

It is interesting to notice the peak of local galaxies with \( B - R = 1.5 \) and the spread and shift to the red of the faint objects in the bottom panel. These magnitudes are not k-corrected, which could well explain the relative redening of the faint objects.

5.3 The local CfA2 Luminosity Function (LF)

Zwicky magnitudes have been used to estimate the LF in the CfA2 redshift catalogue (Marzke et al. 1994). Marzke et al. used a Monte Carlo method to estimate how the LF parameters would change if the dispersion \( \sigma_M \) were 0.65 mag (closer to what we find here than the nominal 0.3 mag they used). They found that the true \( M \) should be about 0.6 mag brighter, and that a similar conclusion would be reached for a small scale error. Marzke et al. used this estimate to conclude that a combination of incompleteness and a small (0.2 mag) scale error would be sufficient to move the CfA2-North values to those found from CfA2-South.
A detailed analysis of the implications of our new Zwicky magnitude calibration on the LF would require the redshift information and this is left for future work. It is nevertheless possible to use the mean relation that we found to show how we expect the LF to change.

In practice, when fitting the Schechter parameters in equation (1) to observational data, the value of $\phi_*$ is correlated with the values of $M_*$ and $\alpha$. In our case we do not use a fit to data, but just model how the LF changes with an homogeneous linear shift in the magnitude scale. A zero point shift $Z_0$ in the magnitude scale, as in equation (17), will just change the value of $M_*:

$$M_*(\text{corrected}) = M_* + Z_0.$$ (19)

As the LF measures the number density of galaxies per magnitude interval, a linear change in the magnitude scale will shift the amplitude of the LF proportionally to the shift in the magnitude interval, i.e. by $\lambda$ in equation (18). Thus we have:

$$\phi_*(\text{corrected}) = \lambda \phi_*.$$ (20)

Thus the scale and zero point error in the Zwicky system give the following corrections for the CfA2 South results of Marzke et al. (1994)

$$\phi_*(\text{corrected}) = \lambda \phi_* \approx 0.0124 \pm 0.006 h^3 \text{Mpc}^{-3}$$ (21)

$$M_*(\text{corrected}) = M_* + Z_0 \approx -19.3 \pm 0.3.$$ (22)

If applied to the LF for the whole CfA2 survey:

$$\phi_*(\text{corrected}) = \lambda \phi_* \approx 0.025 \pm 0.009 h^3 \text{Mpc}^{-3}$$ (23)

$$M_*(\text{corrected}) = M_* + Z_0 \approx -19.1 \pm 0.2,$$ (24)

where we have added the errors in quadrature. Thus the corrected values are now closer the SAPM and LCRS (see Table 1). This can also be seen in Figure 17, where we compare the corrected CfA2 luminosity function given above with the one corresponding to the SAPM. Note that the later is in the APM $b_J$ band, so that there could be some additional zero-point shifts between them (see §2).

6 CONCLUSION

In this paper we have presented CCD magnitudes for galaxies around 204 Zwicky galaxies which sample over 400 square degrees, extending 6 hours in right ascension. This subsample is drawn entirely from volume V of the Zwicky catalogue. We find evidence for a significant scale error as pointed out by Bothun & Schommer (1982) and Giovanelli & Haynes (1984). This is found by direct likelihood analysis, corrected for Malmquist bias, and also by noticing the angular correlations between the errors (Fig. 13) or the long tail of small objects in the frequency distributions of sizes (Fig. 13). The mean scale magnitude error is quite...
large: $\Delta m_Z \simeq 0.62 \Delta B_{CCD}$, i.e. an error of 0.38 mag per magnitude, while the mean zero point is about $-0.35$ (equation 7), in agreement with Efstathiou et al. (1988).

An illustration of the scale error in the Zwicky system can be seen in the galaxies shown in Figures 12 and 14. These illustrations are strongly suggestive of an observer-bias effect, whereby some fainter galaxies are included in the sample due to their proximity to brighter objects, e.g. object #9 in Figure 12. This bias can also be seen in statistical terms as a long tail of small objects in the frequency distribution of Fig.13.

Huchra (1976) found only a 0.08 mag per magnitude scale error in a calibration of Zwicky galaxies with a photoelectric photometry of 181 sample of Markarian galaxies, which are preferentially spirals. As spiral galaxies are bad tracers of clusters or groups, it is unlikely that these Markarian galaxies include any of the fainter galaxies that contribute to the proximity observer-bias mentioned above (which are mostly ellipticals). This effect would be hard to notice with photoelectric photometry which samples only one object at a time. Note that our analysis is restricted to a narrow range of magnitudes $13.5 < m_Z < 15.5$, as compared to the wider range in Huchra (1976), but this narrow range contains the majority of the galaxies with $m_Z < 15.5$ and therefore dominates all the relevant statistical properties (such as the luminosity function).

Bothun & Cornell (1990) have studied the calibration of the Zwicky magnitude scale using a sample of 107 spiral galaxies. They suggest that the errors in $m_Z$ are minimized if $m_Z$ is an isophotal magnitude at $B = 26.0$ mag arcsec$^{-2}$, although it is clear from their Figure 2 that even within such a small sample of objects there is a 5 mag range of isophotes which give isophotal magnitudes corresponding to $m_Z$. This suggests that our isophotal detection limit should be optimal for this comparison.

Takamiya et al. (1995), also with photoelectric photometry, find evidence for a large scale shift between Volume I and Volumes II and V of the Zwicky catalogue. This shift appears to be of order 0.5 mag mag$^{-1}$ over the range $14 < m_Z < 15.7$, which is comparable to our findings for Volume V, although they find very little effect for Volumes II and V. However, we note that Figure 4 of Takamiya et al. shows $B - m_Z$ vs. $B$, rather than $m_Z$. A similar representation of Figure 8 of this paper looks very similar to Figure 4a of Takamiya et al., which is reasonable, given that Figure 8 implies a strong compression of the $m_Z$ axis, which effectively hides the scale error. Unfortunately, the overlap between the 204 objects in our sample and the 155 objects in the Takamiya et al. data is too small to draw any detailed comparison, given that there are ~600 Zwicky galaxies within the region of our survey.

We have also estimated how this scale error could change the CfA2 luminosity function in [5, 6]. Figure 7 shows how the corrected estimation is now closer to other local estimates, such as the SAPM. Finally, in [8] we give properties of the galaxies in our sample: colours, sizes and ellipticities, providing one of the largest samples of this kind. The local colour frequency distribution can be well approximated by a Gaussian distribution with mean $B - R = 1.45$ and rms deviation of $\sigma = 0.40$. These colours compare well with synthetic $B - R$ colours presented Fukugita et al (1995). But there is a significant number of galaxies with $B - R > 2$ which are preferably found in clusters and groups, while most of the bluer objects with $B - R < 1$ (spirals to irregulars) seem more isolated. These local properties are interesting in the context of galaxy evolution and star formation rates. This will be studied in more detail in future work.

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