Reversal of infall in SgrB2(M) revealed by Herschel*/HIFI observations of HCN lines at THz frequencies

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ABSTRACT

Aims. To investigate the accretion and feedback processes in massive star formation, we analyze the shapes of emission lines from hot molecular cores, whose asymmetries trace infall and expansion motions.

Methods. The high-mass star forming region SgrB2(M) was observed with Herschel/HIFI (HEXOS key project) in various lines of HCN and its isotopologues, complemented by APEX data. The observations are compared to spherically symmetric, centrally heated models with density power-law gradient and different velocity fields (infall or infall+expansion), using the radiative transfer code RATRAN.

Results. The HCN line profiles are asymmetric, with the emission peak shifting from blue to red with increasing J and decreasing line opacity (HCN to H13(CN)). This is most evident in the HCN 12–11 line at 1062 GHz. These line shapes are reproduced by a model whose velocity field changes from infall in the outer part to expansion in the inner part.

Conclusions. The qualitative reproduction of the HCN lines suggests that infall dominates in the colder, outer regions, but expansion dominates in the warmer, inner regions. We are thus witnessing the onset of feedback in massive star formation, starting to reverse the infall and finally disrupting the whole molecular cloud. To obtain our result, the THz lines uniquely covered by HIFI were critically important.

Key words. Submillimeter: ISM – ISM: molecules – ISM: structure – ISM: individual objects: SgrB2(M) – ISM: kinematics and dynamics – Stars: formation

1. Introduction

Massive stars form deep inside molecular clouds. While this obscures the process at infrared or shorter wavelengths, it offers the possibility of investigating high-mass star formation using dust emission and molecular lines from the radio to the far-infrared. The shapes of molecular lines are a particularly powerful tool, since they trace the density and temperature gradients as well as the velocity field. Even when the cores are not spatially resolved, radiative transfer modeling of a variety of lines can constrain their structure. The unique spectral coverage of HIFI opens up the full potential of this technique, since both the excitation of rotational transitions and the dust opacity increase with frequency. While the former allows us to study higher temperatures and densities, the latter leads to stronger continuum radiation, thereby facilitating the production of absorption features.

A particularly important stage of high-mass star formation is the hot core phase. When massive stars begin burning hydrogen, they are still accreting and surrounded by dense molecular gas, which is heated and emits a rich spectrum of molecular lines. Various feedback mechanisms (radiation pressure, pressure increase due to ionization, stellar winds, bipolar outflows) push the gas outwards. Finally, this dominates over gravitational accretion, and the molecular cloud is disrupted.

The giant molecular cloud Sagittarius B2 is the most massive and active star-forming region in our Galaxy. It is located close to the Galactic center at a distance of 7.8 kpc (Reid et al., 2009). SgrB2(Main) is a tight cluster of ultracompact Hii regions (Gaume & Claussen, 1990), associated with a massive hot molecular core. Its luminosity is estimated to be 6.3 × 106 L☉ (Goldsmith et al., 1992), the most luminous hot core of the Galaxy. Both infall and a decelerating outflow were seen by Qin et al. (2008).

In this paper, we present new results for the velocity field of SgrB2(M), based on Herschel/HIFI observations of HCN and complemented by APEX data.

2. Observations

Herschel is a new space observatory with a 3.5 m antenna operating at submm to far-infrared wavelengths (Pilbratt et al., 2010). Its spectral line receiver is HIFI (Heterodyne Instrument for the Far Infrared, de Graauw et al., 2010). The guaranteed time key project HEXOS (Herschel/HIFI Observations of EXtraOrdinary Sources, Bergin et al., 2010) includes a full line survey of

* Herschel is an ESA space observatory with science instruments provided by European-led Principal Investigator consortia and with important participation from NASA.
Fig. 1. Herschel/HIFI observations of HCN in SgrB2(M). The green model has a constant infall, the red model has infall in the outer part and expansion in the inner region. Note the different asymmetries in the lines, which are reproduced by the red model. Upper state energies are stated above the plot.

SgrB2(M) ($\alpha_{2000} = 17^h 47^m 20.35^s$ and $\delta_{2000} = -28^\circ 23' 03.0''$). In March 2010, this was started with HIFI bands 1a, 1b, 2a, and 4b, providing coverage of the frequency ranges 479 – 726 GHz and 1051 – 1121 GHz. Each band contains one rotational transition from HCN, H$^{13}$CN, and HC$^{15}$N (Fig. 1), and two rotational transitions within their first vibrational state, which lies around 1000 K higher in energy (Fig. 2). Transition frequencies of HCN (Thorwirth et al., 2003; Fuchs et al., 2004) are taken from CDMS (Müller et al., 2001, 2005). HIFI spectral scans are carried out in dual beam switch (DBS) mode, where the DBS reference beams lie approximately 6' apart. The wide band spectrometer (WBS) is used as a back-end, providing a spectral resolution of 1.1 MHz over a 4-GHz-wide intermediate frequency (IF) band. To allow separation of the two sidebands, bands 1a, 1b, 2a, and 4b were scanned with a redundancy of 4, 8, 8, and 4, respectively (Comito & Schilke, 2003). The data were calibrated using the standard pipeline released with version 2.9 of HIPE (Ott, 2010), and subsequently exported to CLASS using the HiClass task within HIPE. Deconvolution of the DSB data into single-sideband (SSB) format was performed using CLASS. All the HIFI data presented here, spectral features and continuum emission, are deconvolved SSB spectra. The horizontal and vertical polarizations differ by around 10%, and were averaged. The intensity scale is main-beam temperature, and results from applying a beam efficiency correction of about 0.68.

APEX (Atacama Pathfinder Experiment) is a 12-m telescope located at 5100 m altitude in the Atacama Desert, Chile, one of the optimal sites for submm observations on Earth (Güsten et al., 2006). The lines presented here (Fig. 3) are from a study of the structure of 12 hot cores (Rolffs et al. 2010, in prep.). Since the baseline levels in the ground-based data are not determined reliably enough to extract the continuum, the 850µm continuum is taken from ATLASGAL (Schuller et al., 2009).

3. Modeling

The spherical Monte Carlo radiative transfer code RATRAN (Hogerheijde & van der Tak, 2000) was employed to compute the molecular lines and the dust continuum. Using an iterative by-eye comparison, the fit to the data was obtained in two steps. First, the physical structure (density and temperature) was adapted to match the continuum. Second, the HCN abundance and the velocity were optimized for the lines.

The model we present here is centrally heated, its density following a radial power law with index 1.5. The dust opacity was taken from Ossenkopf & Henning (1994), without either ice mantles or coagulation. At the inner radius, 1340 AU, the density is $1.8 \times 10^8$ cm$^{-3}$ and the temperature is 1500 K. The temperature was computed in an approximate way from central heating, using the diffusion equation in the inner part and balance between heating and cooling in the outer part, assuming an effective temperature for the dust photosphere of 58 K. The temperature has a steeper gradient in the inner part, and falls off to 20 K at the...
outer radius of 5.6 pc, where the density is $7 \times 10^3$ cm$^{-3}$. The model reproduces the continuum very well (Table 1). Its central dust optical depth is 0.4 at 345 GHz, 0.8 at 500 GHz, 1.5 at 700 GHz, and 3.7 at 1.1 THz.

To compute the lines, we used a turbulent 1/e half width of 7 km/s (12 km/s FWHM). The lines were additionally broadened by their high optical depth. We found no evidence of significant variations in the turbulent width with radius. The HCN abundance increases with temperature from $10^{-8}$ at temperatures below 100 K, to $3 \times 10^{-8}$ between 100 and 300 K, to $3 \times 10^{-7}$ at temperatures above 300 K. The first jump is mainly needed to reproduce the intensities of the isotopologues and of vibrationally excited HCN, the second jump serves to fit vibrationally excited $^{13}$CN. These jumps could be due to evaporation of ice mantles or increased chemical production. We used a $^{12}$C/$^{13}$C ratio of 20 and $^{14}$N/$^{15}$N of 600 (Wilson & Rood 1994). The green model has a constant infall velocity of 2 km/s, and the red model has the same velocity in its outer parts, but an expansion in its inner parts. Figure 5 shows the velocity field. We note that the exact shape of this, particularly the transition from infall to outflow, is not constrained by the observations, hence modeling the true velocity field by adopting this spherical approximation is an oversimplification (see Fig. 4).

A comparison with the observations (Figs. 1 and 3) shows that the red model reproduces more closely the observed features. In particular, to reproduce the changes in the asymmetry that occur both from lower to higher J and from higher to lower optical depth (from HCN 6–5 to $^{13}$CN 6–5), a change in the velocity field is required, from infall in the colder, outer parts to expansion in the warmer, inner parts.

### 4. Discussion

We now discuss the velocity field and how it is physically determined.

In a gravitationally contracting sphere with density power-law index $p$, a first guess and upper limit to the infall speed would correspond to free fall, i.e., $v_{ff}(R) = \sqrt{\frac{2GM(R)}{R}}$, where the mass inside the radius $R$ is $M(R) \propto R^{3-p}$ (for $p < 3$), such that $v_{ff} \propto R^{2.5}$. In our model, this free-fall speed is roughly (at a radius of 1 pc) 10 times higher than the 2 km/s of the model. Such high infall velocities can be excluded because they would produce lines that are far more asymmetric than observed.

The mass accretion rate $M = v_{acc}(R)\frac{4\pi R^2\rho}{3}$, so $v_{acc} \propto R^{p-2}$ for a constant mass accretion rate. Hence, infall accelerates in a spherical steady-state model with $p = 1.5$. However, an accelerating infall, at least in the inner part, is inconsistent with the data. For simplicity, we chose a constant infall velocity, corresponding to accretion rates onto the cluster of $3 \times 10^{-7}$ M$_\odot$/yr at
infall model has a constant velocity of \(-2\) km/s. At the transition between infall and expansion, the temperature is 65 K and the density is \(6 \times 10^5\) cm\(^{-3}\).

Fig. 5. Physical structure of the model. Density is plotted as dashed green, temperature as dotted blue. The velocity field of the infall-expansion model is shown as the solid red curve (the infall model has a constant velocity of \(-2\) km/s). At the transition between infall and expansion, the temperature is 65 K and the density is \(6 \times 10^5\) cm\(^{-3}\). The real velocity field must be non-spherical, and infall could well continue in the inner part (see Fig. 4 and also In et al. (2008)). The peak on top of the HCN 12–11 absorption (Fig. 4) could be a hint of that infall. Since the H\(^1\)CN transitions are less asymmetric than modeled points in that direction, one may suspect that a linear combination of the pure infall and expanding models would produce the line spectra very well. For the HCN lines, because of their high optical depths, such a linear combination of intensities does not naturally reproduce the observed spectra; more sophisticated modeling, in which we abandon the spherical approximation, would be required. For overlying model and data (Figs. 1, 2, and 3), we assumed a source density of \(10^{-1}\) M\(_\odot\)/yr at a radius of \(10^5\) AU. The excitation rate reaches \(8 \times 10^{-2}\) M\(_\odot\)/yr at \(2 \times 10^4\) AU.

The changing asymmetry of the HCN lines clearly indicates a reversal of infall. Expansion motions dominate in the inner, warmer parts, and infall dominates in the outer, colder parts of the core. This can be naturally explained by the onset of feedback from massive stars. Our radiative transfer modeling demonstrates the power of constraining the source structure by fitting line shapes. A model that successfully reproduces all the features cannot have spherical symmetry, and the exact geometry probably has to be constrained by interferometric observations, to determine the free parameters. Although the Herschel/HIFI data do not have the desired spatial resolution, they place strong constraints on any such model, since they have to be reproduced. We propose a scenario in which the cloud gravitationally contracts at significantly below the free-fall speed. A star cluster forms by fragmentation and accretion. The stars, especially the massive stars that dominate the luminosity, provide various feedback mechanisms to counteract the contraction. The additional pressure, be it thermal, radiative, or turbulent, decelerates the contraction, then dominates first in the inner part, before finally disrupting the whole cloud. SgrB2(M) is just beginning to drive out the gas, while large-scale global infall and probably also infall among very localized pathways in the interior remains ongoing.

This study illustrates the importance of high-frequency lines in constraining the source structure, and demonstrates the great potential of HIFI, which delivers velocity-resolved spectra at these frequencies. Ground-based telescopes may not be able to reach high enough frequencies. In our case, detecting the HCN 12–11 line was necessary in order to unambiguously trace the reversal of infall, so this result is unique to HIFI. We expect that a systematic study of high-J HCN transitions, even higher than the 12–11 line seen here, with a high signal-to-noise ratio, and toward a sample of hot core sources, would be highly rewarding in finding more sources in this particular stage of evolution.

5. Conclusions

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References

Bergin, E., Phillips, T., Comito, C., et al. 2010, A&A, this volume

Comito, C. & Schilke, P. 2002, A&A, 395, 357

de Graauw, T. et al. 2009, A&A, this volume

Fuchs, U., Bruenken, S., Fuchs, G. W., et al. 2004, Z. Naturforsch. A, 59, 861

Gaume, R. A. & Clausen, M. J. 1990, ApJ, 351, 538

Goldsmith, P. F., Lis, D. C., Lester, D. F., & Harvey, P. M. 1992, ApJ, 389, 338

Gusten, R., Nyman, L. A., Schilke, P., et al. 2006, A&A, 454, L13

Hogerheide, M. R. & van der Tak, F. F. S. 2000, A&A, 362, 697

Krumholz, M. R. & Matzner, C. D. 2009, ApJ, 703, 1352

Müller, H. S. P., Schöder, F., Stutzki, J., & Winnewisser, G. 2005, J. Mol. Struct., 742, 215

Müller, H. S. P., Thorwirth, S., Roth, D. A., & Winnewisser, G. 2001, A&A, 370, 149

Ossenkopf, V. & Henning, T. 1994, A&A, 291, 943

Ott, S. 2010, ASP Conference Series, Astronomical Data Analysis Software and Systems XIX, Y. Mizumoto, K.-J. Morita, and M. Ohishi, eds., in press

Pilbratt, G., Riedinger, J., Passvogel, T., et al. 2010, A&A, in press

Qin, S., Zhao, J., Moran, J. M., et al. 2008, ApJ, 677, 353

Reid, M. J., Menten, K. M., Zheng, X. W., Brunthaler, A., & Xu, Y. 2009, ApJ, 705, 1548
