Properties and nature of Be stars  

29. Orbital and long-term spectral variations of γ Cassiopeæ

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Release November 17, 2011

ABSTRACT

A detailed analysis of more than 800 electronic high-resolution spectra of gamma Cas, which were obtained during a time interval of over 6000 days (16.84 yrs) at several observatories, documents the smooth variations in the density and/or extend of its circumstellar envelope. We found a clear anticorrelation between the peak intensity and FWHM of the Hα emission, which seems to agree with recent models of such emission lines. The main result of this study is a confirmation of the binary nature of the object, determination of a reliable linear ephemeris $T_{\text{min}, RV} = \text{HJD} (2452081.9 \pm 0.6) + (203.52 \pm 0.08) \times E$, and a rather definitive set of orbital elements. We clearly demonstrated that the orbit is circular within the limits of our measurements and has a semi-amplitude of radial-velocity curve of $4.30 \pm 0.09$ km s$^{-1}$. No trace of the low-mass secondary was found. The time distribution of our spectra does not allow a reliable investigation of rapid spectral variations, which are undoubtedly present in the spectra. We postpone this investigation for a future study, based on series of dedicated whole-night spectral observations.

Key words. stars: early-type – stars: binaries – stars: Be – stars: individual: gamma Cas

1. Introduction

The well–known Be star of spectral type B0IVe γ Cassiopeæ (27 Cas, HR 264, HD 5394, HIP 4427, MWC 9, ADS 782A), is one of the first two Be stars ever discovered (see Sechter 1966) and a member of a visual multiple system. It exhibits spectral and brightness variations on several timescales. It underwent two major shell phases in 1935–36 and 1939–40. Afterwards, it appeared briefly as a normal B star. Emission strength of the Balmer lines and the brightness of the star in the visual region had been rising slowly during the rest of the 20th century. The observational history of the star has been summarized in detail by Harmanec (2002).

In 1976, γ Cas was identified as an X-ray source. This discovery started a long debate over whether the source of X-rays is the star itself or whether γ Cas is an X-ray binary with a mass accreting compact binary companion. In an effort to prove the duplicity of γ Cas, Cowley et al. (1976) measured radial velocities (RVs hereafter) on a rich collection of photographic spectra obtained in the years 1941–1967. They were unable to find any RV changes exceeding 20 km s$^{-1}$ or to detect any coherent periodicities between 25 and 40000 km s$^{-1}$. Jarad (1987) measured RVs on 81 medium-dispersion (30 Å mm$^{-1}$) photographic spectra using the cross-correlation technique. They concluded that the RVs vary with a period of $P = 1.21277$ km s$^{-1}$, and a well-defined phase curve. Combining their RVs with those measured by Cowley et al. (1976), they found a period of $P = 0.70516$ km s$^{-1}$, which they interpreted as the rotational period of the star. Robinson & Smith (2000) published a detailed study of the X-ray flux of γ Cas. They found that the X-ray flux varied with a period $P = 1.21277$, which they tentatively interpreted as the rotational period of γ Cas. They used this finding as one of the arguments against the binary scenario for the X-ray flux. More recently, Smith et al. (2006) have reported a coherent periodicity of $P = 0.21581 \pm 0.00004$ from the 1998–2006 optical photometry, prewhitened for variations on longer time scales. The latter authors pointed out that the initial period estimate of $P = 1.21277$ was probably an alias of the correct period near $P = 0.21581$.

Harmanec et al. (2000) measured RVs of the steep Hα emission wings in a series of 295 Ondřejov Reticon spectra spanning nearly 2500 days from 1993 to 2000. After removing the long-term RV changes, they discovered periodic RV variations with a period $P = 203.59 \pm 0.29$, semi-amplitude $K_1 = 4.68$ km s$^{-1}$, and eccentricity $e = 0.26$, which they interpreted as the binary motion around a common centre of gravity with a low-mass companion. They demonstrate that the published RVs from the photographic spectra can also be reconciled with the 203.59 period and discuss
the possible properties of the system. Their result was confirmed by Miroshnichenko et al. (2002), who also measured the RVs of the Hα emission wings in a series of 130 electronic echelle spectra, secured with the 1 m reflector of the Ritter Observatory between 1993 and 2002. These two studies differ in the technique of RV measurements. While Harmanec et al. (2000) measured the RVs manually, sliding the direct and reversed continuum-normalized line profiles within a range of intensities on the computer screen until the best match was obtained, Miroshnichenko et al. (2002) also matched the original and reversed profiles, but used an automatic procedure. They arrived at a period of 205.5±0.38, semi-amplitude of 3.80±0.12 km s⁻¹, and a circular orbit. They discuss several possible reasons why their results differ significantly from those of Harmanec et al. (2000) and Miroshnichenko et al. (2002) also document the cyclic long-term spectral variations of γ Cas over a time interval from 1973 to 2002.

To shed more light on the differences between these two RV studies, to obtain truly reliable orbital elements of γ Cas, and to exclude possible 1 d aliases of the 204 d period, we combined our efforts and analysed the two sets of spectra, complemented by more recent observations from Ondřejov and additional spectra from the Dominon Astrophysical (DAO), Haute Provence (OHP) and Castanet-Tolosan Observatories. The RVs in all these spectra were measured by both measuring techniques – alternatively used by Harmanec et al. (2000) and Miroshnichenko et al. (2002) – and analysed in several different ways. Here we present the results of our investigation. We also studied the long-term and rapid spectral variations of γ Cas in our data.

2. Spectral variations

We have collected and analysed series of electronic spectra from five observatories. They cover the time interval from 1993 to 2010. The journal of the observations is in Table 1, where the wavelength range, time interval covered, the number of spectra, and the spectral lines are given. For more details on the individual datasets, readers are referred to Appendix A.

Table 1. Journal of spectral observations.

| Origin of spectra | Δλ (Å) | ΔT (RJD) | Lines | N |
|-------------------|--------|----------|-------|---|
| Ond | 6300–6730 | 49279–55398 | Ha, He, Si | 439 |
| Rit | 6528–6595 | 49272–52671 | Ha, He, Si | 204 |
| OHP | 4000–6800 | 50414–52871 | Ha, He, Si | 34 |
| DAO | 6155–6755 | 52439–54911 | Ha, He, Si | 136 |
| OHP/CMt | various | 53997–55422 | Ha, He, Si | 13 |

Notes. Ond = 2 m reflector of the Astronomical Institute AS CR Ondřejov, Rit = 1 m reflector of the Ritter Observatory of the University of Toledo, DAO = 1.22 m reflector of the Dominon Astrophysical Observatory, OHP = 1.52 m reflector of the Haute Provence Observatory, Cast = Castanet-Tolosan; Δλ = the wavelength region covered, ΔT = the time interval spanned by each dataset, where times are given in the reduced Julian dates RJD = HJD–2400000, Lines: Ha = Hα, He = HeI 6678 Å, Si = Si II 6347 Å, and Si II 6371 Å.

We fitted the Hα line profiles with a Gaussian profile to obtain their peak intensity (I₀ hereafter) and full width at half maximum (FWHM hereafter). This procedure naturally returns a value of I₀, which is slightly less than the very maximum of the emission profile, which is affected by both the V/R varia-

![Fig. 1. Recent evolution of the Hα line profile. Mid-exposure times of the displayed spectra are in RJD = HJD–2400000.](image)
tions and blending with the neighbouring telluric lines. We do believe, however, that the fitted Gaussian provides an objective measure of the gradual changes in the emission-line strength. We omitted the saturated H\textalpha{} line profiles, of course. One example of a Gaussian fit is in Fig. 5 to show where the FWHM and \textit{I}\textsubscript{p} were measured.

![Fig. 2. Recent evolution of the He\textit{i} 6678 Å (left), Si \textit{II} 6347 Å, and Si \textit{II} 6371 Å (right) line profiles. Mid-exposure times of the displayed spectra are in RJD = HJD-2400000. The vertical axis is in the units of the continuum flux.](image)

![Fig. 3. An example of the Gaussian fit to an H\textalpha{} line profile, which also shows the derived quantities \textit{I}\textsubscript{p} and FWHM. Dashed line: the observed H\textalpha{} profile; solid line: the Gaussian fit to it; dotted line: the continuum level.](image)

![Fig. 4. The secular time variations of \textit{I}\textsubscript{p} (upper panel) and FWHM (bottom panel) of the H\textalpha{} line. The time on abscissa is in RJD = HJD-2400000. The open symbols denote measurements from the Ondřejov Reticon detector, capable of distinguishing only 4000 intensity steps, which is why these measurements deviate systematically from the rest. See the text for details.](image)

![Fig. 5. Comparison of the automatically measured H\textalpha{} emission-wing RVs (ordinate = \textit{y}) with those measured manually (abscissa = \textit{x}). The dashed line is a fitted linear function \textit{y} = 0.968 \textit{x}.](image)

3. RV changes

3.1. Long-term and periodic RV changes and the new orbital solutions

Similar to [Harmanec et al. (2000)](Harmanec2000) and [Miroshnichenko et al. (2002)](Miroshnichenko2002), we measured the steep wings of the H\textalpha{} emission in all unsaturated profiles. The reasons the emission wings should provide a good estimate of the true orbital motion of the Be primary around the common centre of gravity with the secondary were recently summarized in detail by Ruzdjak et al. (2009). To them, we can add that [Poeckert & Marlborough (1978)](Poeckert1978) modelled the H\textalpha{} emission of \gamma{} Cas, and their model showed that the H\textalpha{} emission wings originate in regions that are much closer to the star than the radiation that is forming the upper part of the line.

While the CCD detector recognizes 60000 intensity steps. This leads to a systematic difference for strong emission lines.

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\footnote{Although the original Reticon detector and the currently used CCD detector were attached to the same camera of the coudé spectrograph and have the same spectral resolution, the control electronics of the Reticon detector allowed distinguishing only 4000 steps in intensity, which also shows the derived quantities \textit{I}\textsubscript{p} and FWHM. Dashed line: the observed H\textalpha{} profile; solid line: the Gaussian fit to it; dotted line: the continuum level.}
were averaged.

The RVs measured by the first method as manual and those measured by the second method as automatic to distinguish them in following sections. In Fig. 5 the automatic Hα emission RVs are plotted vs. the manually measured ones to see whether there is any systematic difference between the two methods. We fitted the data with a linear relation and somewhat surprisingly the slope was found to be 0.968±0.009; i.e., the automatic method finds a slightly narrower total range of RV variations than the manual one. All individual RV measurements on the steep wings of the Hα emission are published in detail in Table 2 for the manual, and in Table 3 for the automatic measurements.

The Hα emission RVs measured manually and automatically are plotted vs. time in Fig. 6. Additional time plots for other measured features can be found in Appendix A. Figure 6 shows what surprisingly the slope was found to be 0.968±0.009; i.e., the automatic method finds a slightly narrower total range of RV variations than the manual one. All individual RV measurements on the steep wings of the Hα emission are published in detail in Table 2 for the manual, and in Table 3 for the automatic measurements.

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Moreover, we also attempted to measure the RVs of other available spectral lines (He I 6678 Å, Si II 6347 Å and Si II 6371 Å) to see if they undergo similar time changes. We primarily measured also the outer emission wings of these lines, but for the He I 6678 Å line it was possible to obtain a relatively accurate RV of the central absorption core. The RVs measured on the emission wings of the Si II 6347 Å and Si II 6371 Å lines were averaged.

Two methods of RV measurement were used as follows.

1. **Manual measurements** were carried out in the program SPEFO, written by Dr. J. Horn and more recently improved by Dr. P. Škoda and Mr. J. Krpata (see Horn et al. 1992, Škoda 1996). The user can slide the direct and reversed line profile on the computer screen until a perfect match of the selected parts of the profiles is achieved. The advantage of this, admittedly a bit tedious procedure, is that the user actually sees all measured line profiles and can avoid any flaws and blends with the telluric or weak stellar lines. It is invaluable for measurements of weak spectral lines, where any automatic method can be easily fooled by the noise. The same measuring technique has also been used by Harmanec et al. (2000). One of us, JN, also re-measured all Reticon spectra used in their study. Plots of the old vs. new measurements are available as Fig. A3 in Appendix A. Considering the good agreement of both measurements, we used the mean value of the original and new RV measurements for each studied feature. All lines (whenever available) were measured with this technique.

2. **Automatic measurements** were obtained with a program written by AM, which also shifts the direct and reversed line-profile images for a selected range of relative intensities in the continuum units to find a minimum difference between them. The advantages of this method are that it is fast and impersonal. A potential danger is that it can be fooled by flaws and blends in some particular cases. Only the Hα emission wings were measured with this method.

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the range of periods from 3000 d to 50 d, while the periods from 2 to 0.5 are shown in the lower panels. The periodograms are flat, with a close to a value of 1 for all periods between 2 and 50 d, which is why we do not show these parts of the periodograms in the figures. One can see that the combination of RVs from several observatories, which are different from each other in their local times, safely excluded the one-day aliases.

Since no obvious signs of the secondary companion are seen in the spectrum, we adopted y Cas as a single-line spectroscopic binary. We derived a number of orbital solutions for the Hα emission RVs, prewhitened for the long-term changes with HEC13. We used the program SPEL (written by Dr. J. Horn and never published) for this purpose. The program has already been used in several previous studies, e.g., Harmanec (1983, 1984), Koubský et al. (1985), Štefl et al. (1990), and Horn et al. (1992, 1994).

Our first goal was to decide whether the orbital eccentricity found by Harmanec et al. (2000) is real or whether the orbit is actually circular as concluded by Miroshnichenko et al. (2002). In Table 5 the eccentric-orbit solutions for the manually and automatically measured Hα emission-wing RVs are compared. Lucy & Sweeney (1971) have pointed out that observational uncertainties may cause the estimated eccentricity to be biased when the eccentricity is low. The probability that the true eccentricity is zero can be calculated, and this is given in the column "L–S test". If this probability is greater than 0.05 we accept the hypothesis that the true eccentricity is zero at a 5% confidence level.

To shed more light on the problem, we split both manually and automatically measured RVs into a number of data subsets, each of them covering a time interval not longer than three consecutive orbital periods and containing enough observations to define the orbital RV curve. We used the original, not the prewhitened RVs. The phase diagrams for a period $P = 203.52$ d for all selected data subsets are shown in Fig. 9. We derived the elliptical-orbit solutions for them, again testing the reality of the non-zero orbital eccentricity after Lucy & Sweeney (1971). To always find the solution with the smallest rms error, we started the trial solutions for each subset with initial values of $\omega$ for four possible orientations of the orbit, namely $45^\circ$, $135^\circ$, $225^\circ$, and $315^\circ$.

The corresponding orbital elements for the selected subsets, together with the Lucy–Sweeney test, are summarized in Table 5 for manually measured RVs and in Table 6 for automatically measured RVs. The results show very convincingly that the true orbit must be circular (or has a very low eccentricity, which is beyond the accuracy limit of our data). Although the L–S test detected a definite eccentricity for several subsets, the individual values of the longitude of periastron are basically accidental, and that is probably the strongest argument for the eccentric-orbit solutions not being trusted.

### Table 4. Eccentric-orbit solutions based on the Hα emission-wing RVs measured manually and automatically and prewhitened for the long-term changes with the program HEC13.

| Solution No.: | 1) | 2) |
|---------------|----|----|
| Element       | Manual | Automatic |
| $P$ (d)       | $203.36 \pm 0.10$ | $203.08 \pm 0.11$ |
| $T_{RVmin}$ (RJD) | 52083±17 | 52080±13 |
| $\omega$ (°)  | 19\±30 | 147\±22 |
| $e$           | 0.048\±0.027 | 0.072\±0.030 |
| $K_1$ (km s$^{-1}$) | 3.88\±0.10 | 3.93\±0.12 |
| $\gamma$ (km s$^{-1}$) | 0.164\±0.071 | 0.180\±0.081 |
| rms (km s$^{-1}$) | 1.772 | 1.948 |
| L–S test      | 0.183 | 0.064 |
| No. of RVs    | 757 | 700 |
Table 5. The eccentric-orbit solutions for the subsets of manually measured Hα emission-wing RVs, shown as phase plots in Fig.9

| N_i | e   | ω     | K_1   | L–S test | N |
|-----|-----|-------|-------|----------|---|
| 2   | 0.052±0.073 | 141±38 | 4.30±0.24 | 0.734 | 25 |
| 3   | 0.313±0.082 | 165±12 | 5.43±0.44 | 0.001 | 42 |
| 4   | 0.416±0.083 | 108.9±8.3 | 6.46±0.58 | 0.000 | 42 |
| 5   | 0.14±0.14   | 160±41 | 4.37±0.60 | 0.672 | 25 |
| 7   | 0.10±0.11   | 322±47 | 6.88±0.88 | 0.674 | 50 |
| 9   | 0.109±0.067 | 201±36 | 4.10±0.30 | 0.299 | 42 |
| 10  | 0.120±0.040 | 263±16 | 5.20±0.20 | 0.026 | 61 |
| 12  | 0.086±0.055 | 149±28 | 4.64±0.26 | 0.353 | 46 |
| 13  | 0.18±0.10   | 308±29 | 3.56±0.36 | 0.270 | 70 |
| 14  | 0.102±0.040 | 147±22 | 4.31±0.21 | 0.038 | 59 |
| 16  | 0.116±0.067 | 105±31 | 3.87±0.22 | 0.287 | 30 |

Notes. N_i = a number of RV subset (the same numbering being also used in Fig.9), L–S test = probability that the eccentricity found is false, N = number of RVs in the subset.

Table 6. The eccentric-orbit solutions for the subsets of automatically measured Hα emission-wing RVs.

| N_i | e   | ω     | K_1   | L–S test | N |
|-----|-----|-------|-------|----------|---|
| 2   | 0.30±0.14 | 322±24 | 4.37±0.81 | 0.174 | 24 |
| 4   | 0.42±0.11 | 285±17 | 5.82±0.78 | 0.004 | 44 |
| 5   | 0.34±0.10 | 6±20 | 4.74±0.62 | 0.014 | 25 |
| 7   | 0.16±0.12 | 87±49 | 6.0±1.1 | 0.311 | 52 |
| 9   | 0.15±0.19 | 92±53 | 3.38±0.50 | 0.583 | 40 |
| 10  | 0.098±0.050 | 48±23 | 4.12±0.19 | 0.147 | 60 |
| 12  | 0.40±0.21 | 336±9.5 | 5.9±1.2 | 0.054 | 36 |
| 13  | 0.21±0.14 | 82±44 | 3.94±0.53 | 0.157 | 59 |
| 14  | 0.167±0.068 | 111±20 | 4.57±0.33 | 0.026 | 46 |
| 16  | 0.26±0.13 | 227±24 | 4.61±0.56 | 0.201 | 27 |

Notes. N_i = a number of a RV subset (the same numbering being also used in Fig.9), L–S test = probability that the eccentricity found is false, N = number of RVs in the subset.

Adopting the circular orbit from now on, we attempted to remove the long-term RV variations with a different approach. Using the manually measured Hα emission-wing RVs, we treated each data subset spanning no more than three orbital periods (last four subsets) and two orbital periods (remaining subsets) as data coming from separate spectrographs, allowing SPEL to derive individual ‘systemic velocities’ for each subset. This way, the orbital solution was again free of the long-term changes but they were removed in discrete velocity steps. The corresponding orbital solution is in Table 7 and individual velocity levels (γ’s) are shown in the top panel of Fig.6 and listed in Table A.2 in Appendix A. At first sight, this way of removing the long-term changes might seem less accurate, but it leads to smaller rms error for the solution than the removal via HEC13. There are two reasons for that:

1. The HEC13 program computes normal points from RV subsets covering constant time intervals, while for γ velocities the length of subsets was chosen more suitably to our data distribution in time.
2. The HEC13 program computes the normal points as weighted means, but the γ velocities computed with the SPEL program are elements of the Keplerian orbital model.

Fig. 9. Phase diagrams for subsets of Hα emission-wing RVs measured manually. The ordinates of all plots are RVs in km s⁻¹ and the individual subsets vertical axis is in RV and the different ranges reflect the fact that the original, not prewhitened RVs are used. Trial orbital solutions for these subsets are in Tabs. 5 and 6.

For completeness, a solution for an eccentric orbit was also derived but the L–S test gave the probability of 0.18, reassuring us that the eccentricity is spurious. Also for this method of the long-term removal the rms error of the orbital solution is better for the manually than for the automatically measured RVs.

We tested yet another method of removing the long-term variations. It can only be used when one has a RV curve uniformly enough covered by observations. The RVs are averaged over chosen time intervals, and the Hermite polynomials are fitted through these averaged (normal) points. We used the pro-
Table 7. Orbital elements obtained using RVs measured manually (3) and automatically (4) on the Hα line, with the removal of the long-term changes using different γ velocities for individual data subsets.

| Solution No.: | Element | Manual | Automatic |
|---------------|---------|--------|-----------|
| P(d)          | 203.65±0.15 | 203.47±0.14 |
| *T*<sub>min</sub>(RJD) | 52081.42±0.81 | 52081.99±0.95 |
| *K*<sub>1</sub>(km s<sup>-1</sup>) | 4.08±0.10 | 4.14±0.13 |
| rms(km s<sup>-1</sup>) | 1.657 | 1.908 |
| No. of RVs    | 757 | 700 |

Table 8. Orbital elements obtained using the Hα RVs measured manually (5) and automatically (6) and removing the long-term RV variations via a Hermite–polynomial fit through the locally derived γ velocities.

| Solution No.: | Element | Manual | Automatic |
|---------------|---------|--------|-----------|
| P(d)          | 203.52±0.076 | 203.37±0.089 |
| *T*<sub>min</sub>(RJD) | 52081.80±0.62 | 52082.07±0.76 |
| *K*<sub>1</sub>(km s<sup>-1</sup>) | 4.29±0.090 | 4.26±0.11 |
| γ(km s<sup>-1</sup>) | 0.018±0.064 | -0.018±0.075 |
| rms (km s<sup>-1</sup>) | 1.592 | 1.825 |
| No. of RVs    | 757 | 700 |

The net orbital RV curves and the corresponding O–C residua from the orbital solutions are shown in Fig. 10 for the manual and in Fig. 11 for the automatically measured Hα RVs. The rms errors per one observation of the Keplerian fit no. 5 based on manually measured RVs (no. 6 in the case of automatic measurements) are shown as short abcissae in the upper right corners of both figures. The RVs prewhitened for the long-term changes, on which these best solutions 5 and 6 are based, are also presented in Tables 2 and 3.

We did several preliminary tests of the possible rapid variations of γ Cas. Although the results clearly demonstrated their presence, the quantitative results were inconclusive so we decided to postpone analysis of rapid spectral changes for a future study, based on dedicated series of whole-night spectral observations.

4. Interpretation of results

4.1. Spectral variations

We have confirmed the continuing presence of long-term variations in the Hα emission profile using the Gaussian fit and RV measurements. The former procedure gives more objective results than a direct measurement of the peak height and the full width of the line at half maximum (FWHM), because the observed height of the highest peak of a very strong emission line partly reflects the long-term V/R changes, while the measured
FWHM can be affected by the presence of several telluric water vapour-lines of variable strength. Figure 4 shows that the \( I_p \) and FWHM of the Ha emission are anti–correlated over the interval covered by our observations and seem to also show some variability on shorter time scales. We carried out period searches for both these quantities prewhitened for the long-term changes, but no significant periodicities were detected, and trial plots of prewhitened data did not show any evidence of variations connected with the orbital period.

It is probable that the observed long-term changes reflect some variations in the density and/or the extent of the circumstellar disk around the Be primary. Using optical interferometry, Quirrenbach et al. (1997) and Tycner et al. (2006) were able to resolve the envelope around the primary component of \( \gamma \) Cas and estimate that its inclination must be higher than 46° and 55°, respectively. Adopting a reasonable assumption that the rotational axis of the disk is roughly perpendicular to the binary orbit, we can conclude that the binary system is also seen under an orbital inclination higher than some 45°. We then suggest the following possible interpretation of the observed changes.

It is very probable that the envelope around the Be primary is rotationally supported and that its linear rotational velocity decreases with the axial distance from the star. Regardless of the process causing secular variations of the disk, one can assume that the increase in the height of the emission peak reflects the presence of more emission power at lower rotational velocities, thus implying either an increase in the density of the outer parts of the disk and/or an increase in the geometrical extent of the disk. This must naturally decrease the observed FWHM of the emission as observed. This qualitative scenario seems to be supported by the line-profile calculations published by Silaj et al. (2010).

### 4.2. Orbital motion

As mentioned in Sect. 1, one of the principal motivations of this study was to resolve the differences in the orbital solutions obtained by Harmanec et al. (2000) and by Miroshnichenko et al. (2002) and to arrive at a more definitive set of the orbital elements. We were able to combine both independent datasets and complement them with more recent spectra. In Sect. 3, we carried out a number of various tests, analysing separately the RVs measured by a manual and an automatic technique. We also tested the effect of the different ways of data prewhitening on the resulting elements. Special attention was paid to the test of whether the orbit is actually circular or has a significant eccentricity.

The principal results are the following.

1. The binary orbit of \( \gamma \) Cas is circular, at least within the limits of the accuracy of our data, as concluded by Miroshnichenko et al. (2002).

2. The resulting value of the orbital period is now well constrained by the data at hand, and it is robust with respect to different ways of analysis. From all experiments, we were finding values between 203.5 and 203.6, close to the value already found by Harmanec et al. (2000). An inspection of all trial solutions shows that the solutions based on manually measured RVs have rms errors that are systematically lower for \( \approx 15\% \) in comparison to the solutions for the automatically measured Ha emission-wing RVs. We therefore conclude that solution 5 of Table 8 is the best we can offer and suggest the following linear ephemeris for the epoch of RV minimum to be used in the future studies of this binary:

\[
T_{\text{min, RV}} = \text{HJD} (2452081.89 \pm 0.62) + (203.523 \pm 0.0076) \times E.
\]

3. The semiamplitude of the orbital motion is close to 4 km s\(^{-1}\) for all solutions. The recommended value from solution 5 is \( K_1 = 4.297 \pm 0.090 \) km s\(^{-1}\), implying the mass function\( f(M) = 0.00168 M_\odot \). For comparison, Harmanec et al. (2000) and Miroshnichenko et al. (2002) obtained semi- amplitudes of 4.68 \pm 0.25 km s\(^{-1}\) and 3.80 \pm 0.12 km s\(^{-1}\), respectively. If we adopt the inclination value of 45° and the primary star mass \( M_1 = 13 \) M\(_\odot\), suggested by Harmanec et al. (2000), we can estimate the secondary star mass \( M_2 = 0.98 M_\odot\). If the system is at a post mass-transfer phase, then the secondary might be a hot helium star that could be directly detectable in the UV region of the electromagnetic spectrum.

Note: After our paper was accepted for publication, we had the privilege to read a preliminary version of another study of \( \gamma \) Cas kindly communicated to us by Dr. Myron A. Smith and his coauthors. They analysed in particular a smaller and partly independent set of RVs to obtain their own orbital solution. Their results are compatible with our final solution within the respective error bars.

Acknowledgements. We gratefully acknowledge the use of spectrograms of \( \gamma \) Cas from the public archives of the Elodie spectrograph of the Haute Provence Observatory. Our thanks are also due to Drs. P. Chadima, M. Dovčík, P. Hadrava, J. Kubát, P. Mayer, P. Škoda, S. Steffl, M. Wolf, and P. Zasche, who obtained some of the spectra used in this study. Constructive criticism and a careful proofreading of the original version by an anonymous referee helped to improve the paper and are gratefully acknowledged. We thank Dr. M.A. Smith and his collaborators for allowing us to see a preliminary version of their new complex study of \( \gamma \) Cas before it was submitted for publication and for useful comments. This research was supported by the grants 205/06/0304, 205/08/B005, and P209/10/0715 of the Czech Science Foundation, from the Research Programme MSM0021620860 Physical study of objects and processes in the solar system and in astrophysics of the Ministry of Education of the Czech Republic, and from the research project AV0Z10030501 of the Academy of Sciences of the Czech Republic. The research of JN was supported from the grant SVV-263301 of the Charles University of Prague, while PK was supported from the ESA PECS grant 98058. We acknowledge the use of the electronic database from the CDS Strasbourg, and the electronic bibliography maintained by the NASA/ADS system.

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Appendix A: Details of spectroscopic observations and their analyses

A.1. Observational equipment

Here we provide more details on the spectra used in this study (see Table A.1 and their reduction:

1. Ondřejov spectra: All 439 spectra were obtained in the coudé focus of the 2.0 m reflector and have a linear dispersion 17.2 Å mm\(^{-1}\) and a 2–pixel resolution \(R \sim 12600\) (~11–12 km s\(^{-1}\) per pixel). The first 318 spectra were taken with a Reticon 1872RF detector. Complete reductions of these spectrograms were carried out by Mr. Josef Havelka, Mr. Pavol Habuda and by PH in the program SPEFO. The remaining spectra were secured with an STIe–5 800×2000 CCD detector. Their initial reductions (bias subtraction, flatfielding, extraction of 1–D image and wavelength calibration) were done by MS with the IRAF program.

2. DAO spectra: All 136 spectra were obtained in the coudé focus of the 1.22 m Dominion Observatory reflector by SY, who carried out initial reductions (bias subtraction, flatfielding, extraction of 1–D image). The wavelength calibration of the spectra was carried out by JN in SPEFO. The spectra were obtained with the 32121H spectrograph with the IS32R image slicer. The detectors were UBC–1 4096×200 CCD for the data prior to May 2005 and STIe–4 4096×2048 CCD for the data after May 2005. The spectra have a linear dispersion of 10 Å mm\(^{-1}\) and 2–pixel resolution \(R \sim 21700\) (~7 km s\(^{-1}\) per pixel).

3. OHP spectra: Public ELODIE archive\(^5\) of the Haute Provence Observatory (Moultaka et al. 2004) contains 35 echelle spectra obtained at the 1.93 m telescope. They have resolution \(R \sim 42000\). Initial reductions (bias subtraction, flatfielding, extraction of 1–D image, and wavelength calibration) was carried out at the OHP. We only extracted and studied the red parts of the spectra.

4. Ritter spectra: All 204 spectra were obtained with a fiber-fed echelle spectrograph at the 1 m telescope of the Ritter Observatory of the University of Tokyo. We obtained spectra in form of ASCII table covering only region close to H\(\alpha\) spectrogram. The resolution of the spectra is \(R \sim 26000\). Initial reductions of spectra (bias subtraction, flatfielding, extraction of 1–D image and wavelength calibration) were carried out at the Ritter Observatory with the IRAF program.

5. Castanet Tolosan and OHP spectra: We downloaded these spectra from the Be Star Spectra database\(^6\). All of them were obtained by CB with several different spectrograph\(^7\). Only spectra with a resolution comparable to spectra obtained at the rest of observatories were used in the study. Initial reductions (bias subtraction, flatfielding, extraction of 1–D image, and wavelength calibration) were carried out by CB.

For all individual spectrograms, the zero point of the heliocentric wavelength scale was corrected via the RV measurements of selected unblended telluric lines in SPEFO (see Horn et al. 1996, for details).

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Table A.1. Orbital elements obtained using RVs measured on the emission wings of He\(\alpha\) 6678 Å line, absorption core of the He\(\alpha\) 6678 Å line, and emission wings of the Si\(\text{II}\) 6347 Å and Si\(\text{II}\) 6371 Å lines. RJD = HJD–2400000

| spectral line | He\(\alpha\) 6678 Å | He\(\alpha\) 6678 Å | Si\(\text{II}\) 6371 Å |
|--------------|-----------------|-----------------|------------------|
| \(P(d)\)     | 203.52(fixed)   | 203.52(fixed)   | 52096.4±3.5      |
| \(T\text{max}(\text{RJD})\) | 52085.3±3.1     | 52081.4±2.6     | 52096.04±3.5     |
| \(\gamma_{\text{K}}(\text{km s}^{-1})\) | 3.60±0.43       | 5.22±0.50       | 3.14±0.45        |
| \(\gamma_{\text{rms}}(\text{km s}^{-1})\) | 0.16±0.27       | 0.18±0.32       | 0.29±0.28        |
| \(\Delta(\text{km s}^{-1})\) | 6.179           | 7.548           | 6.357            |
| \(N\)         | 560             | 572             | 550              |

A.2. Additional RV measurements

We measured RVs on emission wings and absorption core of He\(\alpha\) 6678 Å and emission wings of Si\(\text{II}\) 6347 Å and Si\(\text{II}\) 6371 Å lines. The program SPEFO was used to the task. The precision of these RV measurements is quite low, since the relative flux in the lines is only several percent greater than in surrounding continuum (see Fig. 2). One could be easily misled during measurements, because measured lines are deformed with continuum fluctuations, and they blend with telluric lines. Despite these complications RVs measured on these lines exhibit long–term variations very similar to the variations that can be seen in Fig. 3. RVs measured on emission wings of He\(\alpha\) 6678 Å line, absorption core of He\(\alpha\) 6678 Å line, and emission wings of Si\(\text{II}\) 6347 Å and Si\(\text{II}\) 6371 Å lines are shown in Fig. A.1.

The long–term variations were removed with the program HEC13, using the 200 d normals and \(\epsilon = 5 \times 10^{-16}\). The model of long-term variations derived by HEC13 is shown in Fig. A.1.

The residuals were searched for periodicity using the HEC27 program. A period near 200 d was detected in all cases, although with a lower significance than the H\(\alpha\) emission RV (see Figs. 4 and 8).

The \(\delta\) statistics periodograms for trial periods from 3000 to 50' are shown in Fig. A.2 where the period \(P = 203^d\) is denoted. The \(\delta\) minimum for this period is not the dominant one only in the case of RV measured on the silicon lines. It is probably due to their low precision and/or incomplete removal of the long-term changes via HEC13.

The circular-orbit solutions were computed for all prewhitened RVs, keeping the orbital period fixed at the value \(P = 203^d\). The corresponding orbital elements are in Table A.1.

Data smoothing with different \(\gamma\) velocities derived with the SPEL program was tested as well. It led to none or a slight (lower than 5%) improvement in rms for the Keplerian fit.

In passing, we wish to mention that the Gaussian fits of the emission RV (see Figs. 7 and 12) the 6 statistics periodograms for trial periods from 3000 to 50' are shown in Fig. A.2 where the period \(P = 203^d\) is denoted. The \(\delta\) minimum for this period is not the dominant one only in the case of RV measured on the silicon lines. It is probably due to their low precision and/or incomplete removal of the long-term changes via HEC13.

The circular-orbit solutions were computed for all prewhitened RVs, keeping the orbital period fixed at the value \(P = 203^d\). The corresponding orbital elements are in Table A.1.

For detailed information on the spectrographs used, see the CB homepage at http://astrosurf.com/buil/.

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\(^5\) URL: http://atlas.obs-hp.fr/elodie/

\(^6\) URL: http://basebe.obspm.fr/basebe/

\(^7\) For detailed information on the spectrographs used, see the CB homepage at http://astrosurf.com/buil/
He I 6678 Å line $a = 1.022 \pm 0.008$. Differences between the measurements of both authors are quite high for He I 6678 Å line emission wings measurements. It is probably because wings of the line are affected by the background noise and because the red peak of the He I 6678 Å line is very low at some point in the V/R cycle.
Fig. A.3. Phase diagrams of RVs prewhitened with HEC13. *Top panel*: the emission wings of the He I 6678 Å line, *Middle panel*: the absorption core of the He I 6678 Å line, *Bottom panel*: the emission wings of the Si II 6347 Å and Si II 6371 Å lines. The period $P = 203.52$ was used.

Fig. A.4. RVs measured by JN plotted vs. RVs measured by PH on spectra obtained at Ondřejov Observatory with Reticon detector. *Top panel*: emission wings of Hα line, *Middle panel*: emission wings of He I 6678 Å line, *Bottom panel*: absorption core of He I 6678 Å line.