INJECTION OF SHORT-LIVED RADIONUCLIDES INTO THE EARLY SOLAR SYSTEM FROM A FAINT SUPERNOVA WITH MIXING FALBACK

A. Takigawa, J. Miki, S. Tachibana, G. R. Huss, N. Tominaga, H. Umeda, and K. Nomoto

ABSTRACT

Several short-lived radionuclides (SLRs), some of which should have formed just prior to or soon after the solar system formation, were present in the early solar system. Stellar nucleosynthesis has been proposed as the mechanism for the production of SLRs in the solar system, but no appropriate stellar source has been found that explains the abundances of all solar system SLRs. In this study, we propose a faint supernova with mixing and fallback as a stellar source of SLRs with mean lives of <5 Myr (26Al, 41Ca, 53Mn, and 60Fe, 107Pd, 129I, and 182Hf) has been inferred from excesses in the abundances of their daughter nuclides in meteorites, which are linearly correlated with the abundance of a parent element (e.g., McKeegan & Davis 2003; Kita et al. 2005). These now-extinct SLRs may provide high-resolution (<0.1 Myr) chronometers for events that occurred during the first several million years of solar system evolution and may also be potential heat sources for asteroidal metamorphism and/or differentiation.

The SLRs with relatively long mean lives, such as 107Pd, 129I, 182Hf, and perhaps 53Mn, may have been products of steady state nucleosynthesis in the Galaxy (Jacobsen 2005), while those with mean lives (τ) of <5 Myr, 10Be (τ = 2.2 Myr), 26Al (τ = 1.03 Myr), 36Cl (τ = 0.43 Myr), 41Ca (τ = 0.15 Myr), 60Fe (τ = 2.2 Myr), and possibly 53Mn (τ = 5.3 Myr), should have been produced either by energetic-particle irradiation in the early solar system or by stellar nucleosynthesis just prior to or shortly after the birth of the solar system. 10Be, which was found in CAIs (calcium-aluminum-rich inclusions; McKeegan et al. 2000), is not produced by stellar nucleosynthesis, but it can be efficiently formed by energetic particle irradiation. On the other hand, 60Fe, the abundance of which in the early solar system also requires a late addition (Tachibana & Huss 2003; Mostefaoui et al. 2005; Tachibana et al. 2006), can only be efficiently produced in stars. Thus, the presence of 10Be and 60Fe in the early solar system suggests that both energetic particle irradiation and stellar nucleosynthesis make contributions to the inventories of solar system SLRs. However, it is not clear yet which process contributed more significantly to the inventory of the SLRs that could be synthesized by both processes, such as 26Al, 36Cl, 41Ca, and 53Mn.

There have been several attempts to find a plausible stellar source or sources for the abundances of SLRs in the early solar system. A low-mass (1.5 Msun) thermonuclear asymptotic-giant-branch (TP-AGB) star cannot produce enough 60Fe to match the initial abundance of 60Fe in the solar system (e.g., Busso et al. 2003; Wasserburg et al. 2006; Sahijpal & Soni 2006). Models for intermediate-mass AGB stars (5 Msun with solar metallicity or 3 Msun with 1/3 solar metallicity) could explain the inferred solar system abundance of 60Fe, as well as the abundances of 26Al, 41Ca, and 107Pd (Wasserburg et al. 2006). However, the probability of encounters between molecular clouds and AGB stars seems to be extremely low (Kastner & Meyers 1994), implying that an AGB star was an unlikely source of SLRs in the solar system. Mass-loss winds from Wolf-Rayet (WR) stars may have been a source of the solar system 26Al, 36Cl, 41Ca, and 107Pd if the nuclides were incorporated into the solar system within a time interval of ~10^5 yr after production (Arnould et al. 2006). However, a Wolf-Rayet (WR) wind alone would not contain enough 60Fe and 53Mn to explain their solar system initial abundances (Arnould et al. 2006).

Type II core-collapse supernovae have also been considered as plausible sources for SLRs. However, most supernova models imply that if a supernova provided 26Al and 41Ca to the solar system, it would also supply 10–100 times more 53Mn than its estimated initial abundance in the solar system (e.g., Goswami & Vanchala 2000; Ouellette et al. 2005; Sahijpal & Soni 2006). This discrepancy could be explained by fallback of the innermost layers, which contain most of the 53Mn, onto a collapsing stellar core (Meyer & Clayton 2000; Meyer 2005). In such a case 53Mn could primarily be derived from a different source, such as the interstellar medium. Wasserburg et al. (2006) proposed that...
53\(^{\text{Mn}}\), and possibly 60\(^{\text{Fe}}\), were injected into the solar system as supernova ejecta with a time interval of \(\lesssim 10^7\) yr after production, long enough for 41\(^{\text{Ca}}\) and 26\(^{\text{Al}}\) to decay completely, and that 26\(^{\text{Al}}\) and 41\(^{\text{Ca}}\) in the solar system may have been produced either by energetic particle irradiation or by an AGB star.

Another problem with supernovae as sources of SLRs is the overproduction of 53\(^{\text{Mn}}\) if all of the 26\(^{\text{Al}}\) in the solar system was derived from supernovae. Although the yield of 53\(^{\text{Mn}}\) depends on the mass loss and initial mass (e.g., Limongi & Chieffi 2006), the expected amount of 53\(^{\text{Mn}}\) injected from a supernova would be, in general, a few times higher than its highest estimate for the early solar system, as we will show later. This problem may still remain even in models for fallback supernovae.

In this study, we propose a supernova with mixing and fallback, with a kinetic energy of explosion slightly less than that for a typical supernova (\(~10^{51}\) erg), as a potential source of 26\(^{\text{Al}}\), 41\(^{\text{Ca}}\), 53\(^{\text{Mn}}\), and 60\(^{\text{Fe}}\) in the early solar system. Faint supernovae such as SN 1997D and SN 1999br have such kinetic energies (e.g., Nomoto et al. 2006). In models for supernovae with mixing fallback, the inner region of the exploding star experiences mixing, some fraction of mixed materials is ejected, and the rest undergoes fallback onto the core (e.g., Umeda & Nomoto 2002, 2005; Nomoto et al. 2006; Tominaga et al. 2007). Nucleosynthesis in a faint supernova with mixing fallback successfully reproduces the elemental abundance patterns of hyper–metal-poor stars (Umeda & Nomoto 2003; Iwamoto et al. 2005; Nomoto et al. 2006).

2. INJECTION OF SUPERNOVA EJECTA INTO THE SOLAR SYSTEM MATERIALS

The abundance of a SLR injected into preexisting solar system materials can be expressed as follows, assuming that injected materials are well mixed with preexisting materials (e.g., Wasserburg et al. 2006; Sahijpal & Soni 2006):

\[
\frac{N_{\text{SLR}}}{N_{\text{SI}}} = \frac{N_{\text{SLR}}^{\text{ejjecta}} f_0 e^{-\Delta \tau}}{N_{0 \text{SLR}}^{\text{SI}} + N_{\text{SI}}^{\text{ejjecta}} f_0},
\]

where \(N_{\text{SLR}}\) and \(N_{\text{SI}}\) are the numbers of the SLR and a stable isotope (SI) for the initial solar system, respectively. Time zero for the solar system is, in practice, considered to be the time of formation of CAIs, the oldest solid materials formed in the solar system. \(N_{\text{SLR}}^{\text{ejjecta}}\) and \(N_{\text{SI}}^{\text{ejjecta}}\) are the numbers of the SLR and SI at the time of their production, respectively. \(N_{0 \text{SLR}}^{\text{SI}}\) is the number of the SI in the preexisting solar system materials. We obtained \(N_{\text{SLR}}^{\text{ejjecta}}\) and \(N_{\text{SI}}^{\text{ejjecta}}\) for supernovae with mixing fallback based on nucleosynthesis models by Nomoto et al. (2006). Their values depend on variable parameters for mixing fallback, such as the scale of the mixing region and an ejection efficiency of mixed materials. Details regarding the determination of \(N_{\text{SLR}}^{\text{ejjecta}}\) and \(N_{\text{SI}}^{\text{ejjecta}}\) in the context of a mixing–fallback model are shown below. \(N_{\text{SLR}}^{\text{ejjecta}}\) and \(N_{\text{SI}}^{\text{ejjecta}}\) were also obtained for non-fallback and fallback supernovae using nucleosynthesis models by Woosley & Weaver (1995), Rauscher et al. (2002), Chieffi & Limongi (2004), and Nomoto et al. (2006) for comparison. \(N_{0 \text{SI}}^{\text{SI}}\) was taken from Anders & Grevesse (1989). The variables \(f_0\) and \(\Delta\) are free parameters in the injection model that represent a dilution factor for supernova ejecta mixed with the preexisting materials and the time interval between the production of the SLR and the CAI formation, respectively.

3. NUCLIDES EJECTED FROM SUPERNOVAE WITH MIXING Fallback

In the standard model for a supernova almost all the materials are ejected, but materials in the very innermost region, which is typically deeper than the incomplete Si-burning layer, fall back onto the star. In the model for a supernova with fallback (Meyer & Clayton 2000; Meyer 2005), such a mass-cut boundary was moved out to the incomplete Si-burning layer or even out to the C/O-burning layer, which resulted in suppressed ejection of 53\(^{\text{Mn}}\).

In the model for a supernova with mixing fallback, we assume two mass-cut boundaries at different depths of a pre-supernova star. The deeper boundary is defined as the initial mass-cut boundary (\(M_{\text{cut}}\)) corresponding to the mass-cut boundary in previous models. The shallower boundary is defined as the outer boundary of the mixing region (\(M_{\text{mix}}\)), which is considered to be the volume between \(M_{\text{cut}}\) and \(M_{\text{mix}}\). Most of materials in the region between \(M_{\text{cut}}\) and \(M_{\text{mix}}\) fall back onto the core due to gravity after complete mixing by Rayleigh-Taylor instabilities, but a small fraction (\(q\)) of homogeneously mixed materials is ejected from the mixing region. Thus, in the present model, supernova ejecta consist of all the nuclides from the region outside of \(M_{\text{mix}}\), and a fraction \(q\) of nuclides within the mixing region between \(M_{\text{mix}}\) and \(M_{\text{mix}}\).

Two mass-cut boundaries (\(M_{\text{cut}}\) and \(M_{\text{mix}}\)) and the fraction of materials ejected from the mixing region (\(q\)) are essential parameters to determine \(N_{\text{SLR}}^{\text{ejjecta}}\) and \(N_{\text{SI}}^{\text{ejjecta}}\) (e.g., Iwamoto et al. 2005; Nomoto et al. 2006). The kinetic energies of explosion of faint supernovae (SN 1997D and SN 1999br) are estimated to be \(~(0.4–0.6) \times 10^{51}\) erg, which is \(~1/3\) to \(~1/2\) times smaller than the typical kinetic energy of supernova explosion. However, an explosion with slightly less kinetic energy does not significantly affect explosive nucleosynthesis (e.g., Woosley & Weaver 1995), and thus we evaluated \(N_{\text{SLR}}^{\text{ejjecta}}\) and \(N_{\text{SI}}^{\text{ejjecta}}\) with different sets of \(M_{\text{cut}}, M_{\text{mix}}, q\) based on a nucleosynthesis model for a solar-metallicity massive star with a kinetic energy of explosion of \(10^{51}\) erg (Fig. 1; Nomoto et al. 2006).

Regarding \(M_{\text{cut}}\), the initial mass cut within the incomplete Si-burning layer suppresses the ejection of 53\(^{\text{Mn}}\) as discussed in fallback supernova models (Meyer & Clayton 2000; Meyer 2005). Although a proper choice of \(M_{\text{cut}}\) in the narrow incomplete Si-burning layer may explain the amounts of the SLRs in the solar system, including 53\(^{\text{Mn}}\), very precise tuning of the model is
required. In the present study, we varied \( M_{\text{cut}} \) from the bottom of the incomplete Si-burning layer to the deeper region as in standard models to check its effect in the context of mixing fallback. We found that the choice of \( M_{\text{cut}} \) affected \( N_{\text{56Fe} \text{ ejecta}} \), but did not significantly change \( N_{\text{56Fe} \text{ ejecta}} \) as long as \( M_{\text{cut}} \) was set deeper than the bottom of the incomplete Si-burning region, where no effective nucleosynthesis of the SLRs discussed in this study takes place. As we will show below, \( f_0 \) is estimated to be much smaller than 1, and the change of \( N_{\text{56Fe} \text{ ejecta}} \) due to change of the initial mass cut has almost no effect on the following discussion. We will show results for \( M_{\text{cut}} \) placed at the boundary of the iron core, where peak temperatures are \((10^{10} \text{--} 8.5) \times 10^9 \text{ K} \).

The position of the outer boundary of the mixing region, \( M_{\text{mix}} \), was varied from within the incomplete Si-burning layer to the bottom of H-burning shell, and the ejection fraction, \( q \), was given values between \( 10^{-4} \) and \( 10^{-2} \). Note that \( M_{\text{mix}} \) and \( q \) were required to be within the C/O-burning region or the He-burning region and \( \sim 10^{-4} \), respectively, to explain the elemental abundances of hyper-metal-poor stars (HE 0107-5240 and HE 1327-2326) by a faint supernova (Iwamoto et al. 2005). The upper limit of \( q \) was determined so that the yield of \( ^{56}\text{Ni} \) did not exceed those for faint supernovae (\( \sim 10^{-3} M_\odot \)) e.g., Nomoto et al. 2006).

4. MODELING THE SOLAR SYSTEM SLRs WITH A SUPERNOVA WITH MIXING FALLBACK

We determined \( f_0 \) and \( \Delta \) to minimize deviations of modeled abundances of \(^{26}\text{Al} \), \(^{41}\text{Ca} \), \(^{53}\text{Mn} \), and \(^{60}\text{Fe} \) (see eq. [1]) from their estimated initial abundances in the early solar system; \(^{26}\text{Al}/^{27}\text{Al} = (5.0 \pm 0.5) \times 10^{-3} \), \(^{41}\text{Ca}/^{40}\text{Ca} = (1.4 \pm 0.14) \times 10^{-8} \), \(^{53}\text{Mn}/^{55}\text{Mn} = (9 \pm 4.5) \times 10^{-6} \), and \(^{60}\text{Fe}/^{56}\text{Fe} = (7.5 \pm 2.5) \times 10^{-7} \) (e.g., Kita et al. 2005). Because the initial abundances of \(^{53}\text{Mn} \) and \(^{60}\text{Fe} \) have not been well determined yet, we chose plausible abundances for them and gave larger uncertainties than for \(^{26}\text{Al} \) and \(^{41}\text{Ca} \) based on ranges of reported initial abundances. Note that the choice of plausible initial abundances for \(^{53}\text{Mn} \) and \(^{60}\text{Fe} \) did not have a significant effect on the following discussion. The uncertainties for the initial abundances were used as weights for minimization of \( f_0 \) and \( \Delta \). Minimization was done in wide varieties of \( f_0 \) and \( \Delta \), of which typical steps were 5% and
100 yr for $f_0$ and $\Delta$, respectively, and we confirmed that the grand minimum was obtained.

5. SHORT-LIVED RADIONUCLIDES FROM SUPERNOVAE WITH MIXING FALLBACK

Calculated initial abundances of $^{26}$Al, $^{41}$Ca, $^{53}$Mn, and $^{60}$Fe, normalized to meteoritic initial abundances for non-fallback supernovae, are shown in Figures 2a–2d as a function of the mass of the exploding star. The dilution factor, $f_0$, is estimated to be in the range of $(0.5 - 6) \times 10^{-4}$, and the typical time interval, $\Delta$, is $\sim 0.7$ Myr. As seen in previous studies, $^{53}$Mn and $^{60}$Fe are, in general, overproduced compared to other SLRs, respectively, irrespective of nucleosynthesis models. Figure 2e shows the case for fallback models based on the nucleosynthesis model of Woosley & Weaver (1995) with no mixing, where the amounts of $^{53}$Mn are much less than that expected in the early solar system (after Meyer & Clayton 2000). As mentioned above, if the mass cut occurs within a narrow incomplete Si-burning layer, it might be possible to explain all the SLRs discussed here with fallback supernovae. However, it requires fine-tuning of the mass-cut region within the layer. Thus the general consequence of fallback supernova models seems to be very little ejection of $^{53}$Mn.

The abundances of SLRs estimated for faint supernova of different masses with mixing fallback are shown in Figure 3 as a function of $M_{\text{mix}}$ and $q$. It can be clearly seen that injection of $^{53}$Mn and $^{60}$Fe is suppressed and that the abundances of SLRs agree with their solar system abundances as long as the outer boundary of the mixing region is located in the C/O-burning layer, where peak temperatures of the shock wave range from $\sim 4 \times 10^8$ to $\sim 4 \times 10^9$ K. On the other hand, the modeled abundances of SLRs cannot explain the solar system values if $M_{\text{mix}}$ is located in the Si-burning or He-burning layers. Iwamoto et al. (2005) proposed that the outer boundary of the mixing layer could be in the upper C/O-burning region or at the bottom of the He-burning layer, which would explain the elemental abundances of hyper–metal-poor stars, HE 0107-5240 and HE 1327-2326, respectively. The former case is consistent with the $M_{\text{mix}}$ required to explain the solar system SLRs. It should be noted here again that the results shown here are not affected by the choice of the initial mass cut ($M_{\text{cut}}$), as long as it occurs in a region deeper than the incomplete Si-burning layer.

In addition to the calculations shown in Figure 3, we calculated the abundances of SLRs with the ejection fraction $q = 10^{-4}$, which has been proposed for hyper–metal-poor stars (Iwamoto et al. 2005). However, the abundance of $^{53}$Mn was about 10 times lower than the estimate for the initial solar system, as in the simple fallback models.

The best estimates for the faint supernovae with mixing fallback are shown in Figure 4, where the calculated abundances of SLRs reproduce their solar system abundances within a factor of 2. It should be emphasized here that a narrow range of mixing-fallback parameters ($M_{\text{mix}}$ and $q$) is not required to match the

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**Fig. 2.** Calculated initial abundances of $^{26}$Al, $^{41}$Ca, $^{53}$Mn, and $^{60}$Fe for supernovae with mixing fallback that have various masses as a function of $M_{\text{mix}}$ and $q$. The calculated abundances are normalized to their estimated initial abundances in the solar system. The dilution factor $f_0$ ranges from $(0.5 - 6) \times 10^{-4}$, and the time interval $\Delta$ ranges from 0.8 to 1.1 Myr. Each panel displays the results of 18–20 sets of calculations; the horizontal axis gives the mass in solar masses of the material inside the outer boundary of the mixing region for each calculation. Dotted and dashed lines show the boundaries between the Si-burning and C/O-burning layers and between the C/O-burning and the He-burning layers, respectively.

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**Fig. 3.** Calculated initial abundances of $^{26}$Al, $^{41}$Ca, $^{53}$Mn, and $^{60}$Fe for supernovae with mixing fallback that have various masses as a function of $M_{\text{mix}}$ and $q$. The calculated abundances are normalized to their estimated initial abundances in the solar system. The dilution factor $f_0$ ranges from $(0.5 - 6) \times 10^{-4}$, and the time interval $\Delta$ ranges from 0.8 to 1.1 Myr. Each panel displays the results of 18–20 sets of calculations; the horizontal axis gives the mass in solar masses of the material inside the outer boundary of the mixing region for each calculation. Dotted and dashed lines show the boundaries between the Si-burning and C/O-burning layers and between the C/O-burning and the He-burning layers, respectively.

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**Fig. 4.** Calculated initial abundances of $^{26}$Al, $^{41}$Ca, $^{53}$Mn, and $^{60}$Fe for supernovae with mixing fallback that have various masses as a function of $M_{\text{mix}}$ and $q$. The calculated abundances are normalized to their estimated initial abundances in the solar system. The dilution factor $f_0$ ranges from $(0.5 - 6) \times 10^{-4}$, and the time interval $\Delta$ ranges from 0.8 to 1.1 Myr. Each panel displays the results of 18–20 sets of calculations; the horizontal axis gives the mass in solar masses of the material inside the outer boundary of the mixing region for each calculation. Dotted and dashed lines show the boundaries between the Si-burning and C/O-burning layers and between the C/O-burning and the He-burning layers, respectively.
solar system abundances of SLRs. The overall agreement of calculated abundances with meteoritic values can be obtained for wide ranges of \( M_{\text{mix}} \) and \( q \) values (Fig. 3), a major advantage of this model compared to the simple fallback supernova models.

6. A FAINT SUPERNova WITH MIXING Fallback AS A Source of solar System SLRs

Previous supernova models, which attempted to explain the initial abundances of solar system SLRs, had the problems of overinjection of \( ^{53}\text{Mn} \) and \( ^{60}\text{Fe} \) compared to \( ^{26}\text{Al} \) and \( ^{41}\text{Ca} \). Such problems can be solved in the faint supernova model with mixing fallback with the following explanations.

\( ^{53}\text{Mn} \) is synthesized mainly in the incomplete Si-burning layer, while \( ^{26}\text{Al} \), \( ^{41}\text{Ca} \), and \( ^{60}\text{Fe} \) are produced more abundantly in outer regions (Fig. 1). Although a small fraction of the \( ^{53}\text{Mn} \) is ejected, most falls back only to the collapsing core when the outer region of the mixing layer is located in the C/O-burning layer. On the other hand, some fractions of other SLRs, produced outside of the mixing layer, are ejected without experiencing the mixing fallback. Together, these features suppress the amount of \( ^{53}\text{Mn} \) injected into the solar system materials compared to other SLRs.

\( ^{26}\text{Al} \) is the only nuclide, among the SLRs considered in this study, that forms abundantly in the He and H layers and the explosive O-burning layer of a massive star with solar metallicity. If the outer boundary of the mixing region is located in the C/O layer, more than 80% of \( ^{26}\text{Al} \) should have its origin in the He and H layers. This is because \( ^{26}\text{Al} \) in the C/O layer was destroyed by the He burning (e.g., Limongi & Chieffi 2006), and most of the \( ^{26}\text{Al} \) in the explosive O-burning layer falls back onto the central remnant. The amount of \( ^{26}\text{Al} \) in the supernova ejecta is thus less affected by the mixing-fallback process, which lowers the \( ^{60}\text{Fe} / ^{26}\text{Al} \) ratio. \( ^{26}\text{Al} \) may be ejected in a mass-loss wind before the supernova explosion, which may explain the possible time interval between \( ^{60}\text{Fe} \) and \( ^{26}\text{Al} \) injection (Bizzarro et al. 2007).

We conclude that \( ^{26}\text{Al}, ^{41}\text{Ca}, ^{53}\text{Mn} \), and \( ^{60}\text{Fe} \) in the early solar system could have been brought from a single massive star that experienced a less-energetic supernova explosion. The estimated dilution factor, \( f_0 \), may provide a geometric constraint on the supernova and the solar system materials as discussed by Sahijpal & Soni (2006). Assuming the spherical symmetric ejection from the supernova, \( f_0 \) can be basically expressed by an injection efficiency of ejecta into the solar system materials (\( \alpha \)) and a solid angle (\( \Omega \)) of the solar system materials (a protosolar molecular cloud core or a proto–solar system disk) on a sphere of ejecta with a radius corresponding to the distance (\( D \)) between the supernova and the solar system materials. If ejecta are injected into the molecular cloud core (0.1 pc in radius, 1 \( M_\odot \) in mass, and an \( \alpha \) value of 0.1; Vanhara & Boss 2002), \( D \) is estimated to be \( \approx 2–5 \) pc. In the case of injection into the proto–solar system disk, as discussed in Ouellette et al. (2005; 100 AU in radius, 0.01 \( M_\odot \) in mass, and an \( \alpha \) value of 1), a \( D \) value of \( \approx 0.3–0.8 \) pc was obtained.

It is difficult to evaluate which case is more plausible for the solar system–forming environment at present. However, in either case, the supernova explosion should have occurred near the solar system materials, which supports the idea that the solar system was born in a star cluster containing massive stars. The lifetime of the star cluster is several to 10 Myr (e.g., Adams & Laughlin 2001), within which only massive stars (>20–25 \( M_\odot \)) explode. The proportion of massive stars in a cluster is low (e.g., Kroupa 2001), and thus it may be highly unlikely that multiple supernovae brought SLRs into the solar system materials within the lifetime of the star cluster. In the model for a faint supernova with mixing fallback, solar system SLRs with mean lives of <5 Myr could be from a single supernova.

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Fig. 4.—Best estimates for the abundances of \( ^{26}\text{Al}, ^{41}\text{Ca}, ^{53}\text{Mn}, \) and \( ^{60}\text{Fe} \) in models for faint supernovae with mixing fallback. The values for the parameters are \( q = 0.001 \), \( M_{\text{mix}} = 3.1 \ M_\odot \) (6.7 \( \times 10^7 \) K; peak temperature of the shock wave), \( f_0 = 1.90 \times 10^{-3} \), and \( D = 1.07 \ Myr \) for 20 \( M_\odot \); \( q = 0.005 \), \( M_{\text{mix}} = 3.0 \ M_\odot \) (1.4 \( \times 10^7 \) K), \( f_0 = 1.34 \times 10^{-4} \), and \( D = 0.83 \ Myr \) for 25 \( M_\odot \); \( q = 0.001 \), \( M_{\text{mix}} = 6.7 \ M_\odot \) (7.8 \( \times 10^7 \) K), \( f_0 = 4.35 \times 10^{-4} \), and \( D = 0.87 \ Myr \) for 30 \( M_\odot \); and \( q = 0.005 \), \( M_{\text{mix}} = 11.8 \ M_\odot \) (5.8 \( \times 10^7 \) K), \( f_0 = 1.64 \times 10^{-4} \), and \( D = 0.75 \ Myr \) for 40 \( M_\odot \). Error bars represent uncertainties of initial abundances of SLRs estimated from meteorites (see text).
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