Carbon-enhanced stars with short orbital and spin periods

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ABSTRACT
Many characteristics of dwarf carbon stars are broadly consistent with a binary origin, including mass transfer from an evolved companion. While the population overall appears to have old-disc or halo kinematics, roughly 2 per cent of these stars exhibit Hα emission, which in low-mass main-sequence stars is generally associated with rotation and relative youth. Its presence in an older population therefore suggests either irradiation or spin-up. This study presents time-series analyses of photometric and radial-velocity data for seven dwarf carbon stars with Hα emission. All are shown to have photometric periods in the range 0.2–5.2 d, and orbital periods of similar length, consistent with tidal synchronisation. It is hypothesised that dwarf carbon stars with emission lines are the result of close-binary evolution, indicating that low-mass, metal-weak or metal-poor stars can accrete substantial material prior to entering a common-envelope phase.

Key words: binaries: general – stars: carbon – stars: activity – stars: chemically peculiar – stars: rotation – white dwarfs

1 INTRODUCTION
Binary-star evolution is responsible for a wide array of astrophysical phenomena, including blue stragglers, Type-Ia supernovae, and gravitational waves (Maeder 1987; Webbink 1984; Abbott et al. 2016). Interactions between binary components may result in deviations from the course of typical single-star evolution, through mass transfer or mergers, leading to changes in stellar structure and atmospheric chemistry (Hurley et al. 2002; Tremmper et al. 2020). A prime example of how binary interactions can influence stellar evolution is given by dwarf carbon (dC) stars. These low-mass, main-sequence stars exhibit spectra analogous to those of evolved carbon stars found at the tip of the asymptotic giant branch (AGB; Dahn et al. 1977), but in advance of dredge-up processes being able to modify surface chemistry.

The peculiar atmospheric chemistry (C/O > 1) in the prototype dC star, G77-61, is thought to originate through mass transfer from an evolved, higher-mass companion (Dearborn et al. 1986). In this picture, as the former primary star ascends the AGB, CNO-processed material is dredged to the surface before stellar winds liberate the consequently carbon-enhanced outer layers (Iben & Renzini 1983). Enriched material is then transferred to the main-sequence companion, polluting its atmosphere and creating a dC star, while the original primary becomes a white dwarf. The exact physical means of mass transfer in these systems remain unknown, but possibilities include classical processes of stellar-wind capture and Roche-lobe overflow, and more-recent models such as wind–Roche-lobe overflow (Abate et al. 2013). Regardless of the detailed mechanisms, metal-poor stars are significantly more susceptible to atmospheric pollution through external carbon enrichment processes (de Kool & Green 1995).

There is little optical evidence for white-dwarf companions to dC stars, where only ≤ 1 per cent of systems exhibit flux consistent with a white dwarf at blue wavelengths (e.g. Heber et al. 1993; Liebert et al. 1994; Green 2013). Nonetheless, recent observational evidence supports a binary formation channel for dC stars. A radial-velocity study of 28 dC stars, each with typically 3–5 measurements obtained over several years, suggests a dC binary fraction of at least 75 per cent (Whitehouse et al. 2018). A larger survey of 240 dC stars, at a lower resolution, also found a high binary fraction, possibly as large as 95 per cent, with a mean orbital period on the order of 1 yr (Roulston et al. 2019).

The first four, well-characterised dC binaries all have orbital periods on the order of years, as established by radial-velocity or astrometric monitoring (Dearborn et al. 1986; Harris et al. 2018): G77-61 (0.67 yr), LSPM0742+4659 (1.23 yr), LP255-12 (3.21 yr), and LP758-43 (11.4 yr). These systems are all single-lined spectroscopic binaries or their astrometric equivalent, with neither optical nor ultraviolet indications of flux from their suspected evolved companions.

The existence of long orbital periods (≥ 100 d) among dC stars is commensurate with better-studied classes of carbon-enhanced binary systems, such as CH, Ba, and carbon-enhanced metal-poor (CEMP) stars, which are all thought to follow analogous evo-
lutionary pathways and which have orbital periods of hundreds to thousands of days (Jorissen et al. 2016; McClure et al. 1980; Lucatello et al. 2005). In contrast, short-period binaries are not necessarily expected, as Roche-lobe overflow tends to be unstable in the case of mass transfer from a higher- to lower-mass star, and the sole population-synthesis study for dC stars verifies that short-period systems fail to result if accretion is inefficient during a common-envelope phase (de Kool & Green 1995).

However, in the same year that three of the long-period orbits were published, the dC star SDSS J125017.90+252427.6 (hereafter J1250) was reported to have a photometric period of 2.9 d, and radial-velocity measurements folded on this period appear to provide a good orbital solution (Margon et al. 2018). This system has two notable features: photometric variability on a period potentially identical to the orbital period, and prominent emission in the Balmer lines and in Ca H & K, all of which are photospheric (Margon et al. 2018). Such emission features are rare among dC stars, with only a handful known among the ~1200 confirmed and suspected candidates in the Sloan Digital Sky Survey (Green 2013). Furthermore, a Chandra study of six emission-line dC stars detected each of them in X-rays – a clear indication of stellar activity (Green et al. 2019).

Motivated by these findings, this study investigates the possibility that emission lines observed in some dC stars arise because of unexpectedly rapid rotation, associated with tidal synchronisation in short-period binaries, or with spin-up due to mass transfer. This paper presents the results of radial-velocity and photometric time-series analyses conducted for seven dC stars with Hα emission, and finds that six stars in the sample are binaries with orbital periods in the range 0.2–4.4 d. All seven sources show photometric variability on periods that are either orbital, or are similarly short. The radial-velocity and photometric observations are summarised in Section 2, the methods and period analyses are discussed in Section 3, the results for each individual target are presented in Section 4, and a discussion follows in Section 5.

## 2 OBSERVATIONS AND DATA REDUCTION

### 2.1 Target selection

The literature was searched to identify dC stars known to exhibit Hα emission (e.g. Lowrance et al. 2003; Green 2013), with six selected for further study based on brightness, visibility, as well as proper motion consistent with the main sequence, and are listed in Table 1. Four have been monitored previously for radial-velocity variations over a several-year baseline with a median of five observing epochs per target (J1015, CLS29, SBSS 1310, CBS 311; Whitehouse et al. 2018). Furthermore, the emission-line dC star J1250 is included in this analysis, though no new data have been obtained for this target.

### 2.2 Radial velocities

Radial-velocity monitoring was conducted using the ISIS spectrograph on the 4.2-m William Herschel Telescope at the Roque de Los Muchachos Observatory. Some of the observations are reported in a previous paper (Whitehouse et al. 2018), and further observing

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**Table 1. Observed dC stars with Hα emission.**

| Target          | J2000 Coordinates (h m s) | (° ′ ″) | SDSS r (AB mag) | Ref. |
|-----------------|--------------------------|--------|----------------|-----|
| SDSS J084259    | 08 42 59.8 +22 57 29     | 17.1   | 1              |
| SDSS J090128    | 09 01 28.3 +32 38 33     | 17.1   | 1              |
| SDSS J110158    | 10 15 48.9 +09 46 49     | 16.9   | 2              |
| CLS 29          | 10 40 06.4 +35 48 02     | 14.9   | 3              |
| SBSS 1310+561   | 13 12 42.4 +55 55 54     | 14.3   | 4              |
| CBS 311         | 15 19 06.0 +50 07 02     | 17.4   | 5              |

Notes: The final column lists either the discovery or the first reference to illuminate the nature of the source: (1) Green 2013; (2) Whitehouse et al. 2018; (3) Mould et al. 1985; (4) Rossi et al. 2011; (5) Liebert et al. 1990. Runs have since been carried out in 2018 Aug, 2018 Feb and Apr, and 2020 Feb. All observations prior to 2020 were carried out using only the blue arm of the spectrograph, without any dichroic, in combination with the R1200B grating. An EEV12 detector was employed without binning, providing wavelength coverage of approximately 5000–6000 Å and a resolving power R ≈ 6700 at the central wavelength (1 arcsec slit). The wavelength range encompasses many strong atomic and molecular lines, including the Mg i triplet, Na i D doublet, and two prominent C2 Swan bands (see Table 2).

In 2020 Feb, both the blue and red arms of the ISIS spectrograph were used, with the GG495 dichroic. The blue arm was equipped with the R600B grating to provide wavelength coverage of 3510–5340 Å, achieving a resolving power R ≈ 2700 at the central wavelength. These data were obtained using the EEV12 detector without binning. The red arm of the spectrograph was employed with the R1200B grating, covering a wavelength range of 5770–6770 Å, where a resolving power of R ≈ 7400 was achieved. These data were obtained using the RED+ detector with no binning. Together, these more recent data cover Hα and the higher Balmer lines.

Beginning in 2019, each star was observed a number of times (typically two to four) during each epoch to probe for short-term changes in radial velocity. Each observation consisted of at least three separate exposures of the science target, with wavelength-calibration frames taken immediately before and after. Individual exposure times for the stars varied between 180 and 600 s typically, for the brighter and fainter sources, respectively. Standard stars and other calibration data were taken at least once per night.

All science exposures were bias subtracted, flat fielded, and extracted using standard methods in IRAF. For extraction, the 2-dimensional spectral trace was established using the standard star observation with the highest signal-to-noise (S/N) of the night and was re-centred for the extraction of each target spectrum. Individual exposures were wavelength calibrated using the appropriate arc frames for each target. Finally, all consecutive target exposures were combined using inverse variance weighting, to create a single spectrum with S/N ≥ 10 for analysis.

### 2.3 Photometry

Time-series photometry used in this study was sourced primarily from the public data releases of the Palomar Transient Factory (PTF), intermediate-Palomar Transient Factory, and Zwicky Transient Facility (ZTF) surveys (Law et al. 2009; Kulkarni 2013; Bellm 2014). The primary science goals of these projects are to monitor the sky for transient events, using the Palomar 48-in Oschin Schmidt
Table 2. Wavelength regions and atomic transitions used for spectral cross-correlation to determine absolute velocities.

| Wavelengths (Å) | Important lines |
|-----------------|-----------------|
| 4092 – 4112     | Hδ             |
| 4330 – 4350     | Hγ             |
| 4851 – 4871     | Hδ             |
| 5163 – 5187     | Mg i b         |
| 5197 – 5215     | Fe i Cr i blends |
| 5258 – 5273     | Fe ii          |
| 5325 – 5331     | Fe i           |
| 5368 – 5376     | Fe i           |
| 5863 – 5924     | Na i D         |
| 6553 – 6573     | Hα             |

Telescope with wide-field imagers, surveying the entire available northern sky roughly every three days. The detection limit of the survey is \( R \approx 20.6 \) mag (Rau et al. 2009), and all data are fully reduced and calibrated by the dedicated facility pipelines (Ofek et al. 2012; Masci et al. 2019).

Data were retrieved for J0901, J1015, and CBS 311 by querying the PTF DR3 and ZTF DR5 databases using their Gaia eDR3 astrometry (Gaia Collaboration et al. 2020). For each target, a cone search was made using a 15-arcsec radius centred on the 2015.5 Gaia positions. Only matches within 0.5 arcsec of the expected position were selected for analysis; this is well within the typical diameter of the point spread function of the surveys. All photometric data with poor quality flags were rejected.

Photometric monitoring of the dC star J0842 was conducted by Kepler during K2 campaign 18, with continuous measurements for two months at a cadence of 30 min (Howell et al. 2014). Other sources of photometry were not used for this target, as the Kepler data are of superior quality.

Finally, CLS 29 and SBSS 1310 were observed with the Transiting Exoplanet Survey Satellite (TESS; Ricker et al. 2015), and are sufficiently bright and isolated that useful data were obtained. CLS 29 was observed in Sector 22, and SBSS 1310 in Sectors 15, 16, and 22 (with an interval of 139 days between Sectors 16 and 22). The duration of each TESS sector is approximately 25 d, with photometric observations taken every 30 min.

### 3 METHODS AND PERIOD ANALYSIS

#### 3.1 Radial-velocity cross-correlation

The most precise measurements of radial-velocity changes are achieved by cross-correlating individual spectra of a given target against its highest S/N counterpart. Such measurements benefit from strong molecular carbon absorption features (e.g. C2 Swan bands) in dC stars. To achieve the best cross-correlation possible, all spectra were first continuum-normalised by fitting a cubic spline. The cross-correlation was performed using the IRAF FXCOR package (Tonry & Davis 1979). This cross-correlation method reduces the velocity error by around 50 per cent compared to a determination of absolute velocity using the few available (weak) atomic lines. These more precise but relative velocity measurements are then used to determine the orbital parameters for each target.

To establish an absolute velocity scale for these data, the highest S/N spectrum of each target was cross-correlated against a continuum-normalised synthetic spectrum template. These spectra were synthesised using AUTOKUR, which is a wrapper for the Kurucz spectral synthesis codes (Kurucz 2005), with the underlying atmospheric structures based on MARCS models (Gustafsson et al. 2008). Stellar parameters were chosen to be representative of low-mass main-sequence stars (i.e. appropriate for late-K to early-M dwarfs), and to reflect the older kinematical ages of the dC population as a whole. Templates were generated for \( T_{\text{eff}} \approx 3500 – 4500 \) K in steps of 250 K, and metallicities \([\text{Fe/H}]\) from \(-2\) to \(0\) in integer steps, with the surface gravity and microturbulence kept fixed at \( \log g (\text{cm s}^{-2}) = 5.0\) and \( v_t = 2.0 \) km s\(^{-1}\), respectively. The synthetic spectra were generated at resolving powers of \( R = 6700\) and 7900 to match the instrumental resolution achieved under standard and best observing conditions at the telescope, respectively.

Because stellar-atmosphere models are not currently available for cool, carbon-rich dwarf stars, the synthetic spectra were generated for stars with solar carbon abundances, at gravities and effective temperatures similar to those expected for a typical dC star. The C2 Swan bands, which contain myriad individual transitions and are useful for relative velocity measurements, are not present in the solar-abundance templates, and would lead to increased measurement errors in velocity if used as the primary source of cross-correlation. Therefore, to minimise the measurement errors, restricted wavelength regions of the synthetic templates – those containing the strongest atomic lines – were selected for cross-correlation. Using these restricted regions significantly improved the quality of each cross-correlation (albeit while still not as precise as the relative velocity measurements), by mitigating spectral mismatches. The wavelength regions used for the synthetic templates are listed in Table 2.

To determine the best synthetic template to measure absolute velocities, each synthetic template was cross-correlated against the highest S/N spectrum of each observed dC star. The template that achieved the strongest cross-correlation function was therefore deemed as the most suitable. In general, templates at cooler temperatures achieved a stronger cross-correlation function than templates at higher temperatures; those with \( T_{\text{eff}} < 4000\) K consistently outperformed their warmer counterparts across the whole metallicity range. Furthermore, templates with metallicity \([\text{Fe/H}] = 0\) typically performed worse when compared to lower metallicity templates. Overall, the resolving power of the templates exhibited no performance trend in cross-correlations. The stellar parameters of the best performing synthetic spectral template were \([\text{Fe/H}] = -1\), \( T_{\text{eff}} = 3750\) K, and \( R = 6700\). This template was therefore adopted to measure absolute velocities, as it consistently possessed strong cross-correlation metrics with high S/N observed spectra.

#### 3.2 Radial-velocity period determinations

The JOKER software is designed to determine orbital periods from radial-velocity data, under the assumption that all variability is the product of two-body gravitational interactions (Price-Whelan et al. 2017, 2020). For a given number of samples drawn from a prior probability distribution function (PDF), it uses the method of rejection sampling to eliminate low-likelihood solutions, generating a random value for each posterior sample, \( r_j \), between 0 and the maximum likelihood in the posterior distribution, \( L_{\text{max}} \). If the likelihood of a specific sample, \( L_j \), does not meet the condition that \( r_j < L_j \), then the sample is rejected. This approach guarantees at least one surviving sample and is particularly useful where radial-velocity data are sparse.

Prior PDFs were selected for orbital eccentricity, velocity semi-amplitude, and systemic velocity as follows. Only circular
orbits were considered, both for simplicity and because the available radial-velocity sampling is insufficient to robustly characterise eccentric orbits. Should a short-period solution result for any given system, then the circular orbit assumption is retrospectively well justified for post-common-envelope binaries (Paczyński 1976; Podsiadlowski 2001). Velocity semi-amplitudes were sampled from a Gaussian PDF; for each target, the mean was set to zero, with the standard deviation set to 110 km s$^{-1}$ (taking the absolute value). The adopted standard deviation is based on the expected velocity semi-amplitude for a binary with a total mass of 1.1 $M_\odot$, a 5 d period, and 60$^\circ$ inclination. The width of the prior PDF on velocity semi-amplitude is intended to avoid biasing the posterior distributions. Finally, systemic velocities were sampled from a Gaussian distribution also centred at zero, with a standard deviation of 100 km s$^{-1}$. Again, this prior PDF is uninformative and should not bias against any actual velocities.

Because the joker analysis is performed on the more precise, relative radial-velocity data, the systemic velocity returned from the analysis is offset from the true value. This was corrected by applying the absolute velocity offsets as determined by the synthetic spectral templates, while propagating the errors in the offset measurement (which are assumed to be Gaussian) with the systemic velocity posterior distribution.

To ensure that an adequate number of posterior samples survive the joker analysis, $10^6$ prior samples are generated for each target, with the maximum posterior sample size set to 8192. However, the number of posterior samples removed during the rejection sampling algorithm in the joker is a function of the maximum likelihood of those posterior samples. Therefore, for targets possessing more constraining data, fewer posterior samples survive.

In the scenario where few samples survive rejection sampling, the posterior distribution returned from the joker is typically unimodal with fewer than 100 samples. Owing to the small number of surviving samples, in order to explore the posterior distribution fully, these surviving samples are used as prior samples to initialise a Markov Chain Monte Carlo (MCMC) simulation. The pymc3 no-u-turn sampler was used with four chains and a tuning length of 1000, yielding a total of 2500 posterior samples per chain (Salvatier et al. 2016; Hoffman & Gelman 2011). This further analysis with MCMC was necessary for the dC stars CLS 29 and SBSS 1310, as only 10 and 74 samples survived rejection sampling, respectively. All other targets possess sufficiently large posterior distributions from the joker such that this additional analysis was omitted.

The orbital parameters for each target are quoted as the median of the final posterior distribution (whether through the joker or MCMC) and presented in Table 3. The median was selected due to the non-Gaussian nature of some of the distributions. The uncertainties in the orbital parameters for each target are quoted as the 16 per cent and 84 per cent confidence limit in the posterior distribution. This analysis was completed for the six targets discussed in this paper, as well as J1250. The radial-velocity curves corresponding to the median orbital parameters derived through this analysis are shown in Figure 1, with the posterior PDFs of each target displayed in Figure A1.

3.3 Periodogram analysis of photometry

All photometric time-series were normalised to the median value for each target, then analysed using the $\delta\chi^2$ implementation of the Lomb-Scargle (LS; Lomb 1976; Scargle 1982) periodogram, and the $\chi^2$ (Huijse et al. 2018) implementation of Multi-Harmonic Analysis of Variance (MHAOV; Schwarzenberg-Czerny 1996). The results of the LS analysis are presented in Figure 2. The LS periodogram is a method to derive the period of oscillation, for a recurrent signal, in unevenly sampled time-series data through the construction of a Fourier-like power spectrum (VanderPlas 2018); it is almost equivalent to performing a least-square fit of a sine wave (cf. Zechmeister & Kürster 2009). Such a periodogram can be relatively difficult to interpret, as the Fourier

![Figure 1](https://example.com/figure1.png)
transform of any continuous signal in the data is convolved with the window function. It can therefore be challenging to identify which periodogram peaks represent real signals in the time-series data. To break the degeneracy between the true and aliased peaks, an LS periodogram was computed for the window function for each target, and the window-function periodogram compared to the periodogram computed from the actual fluxes.

For an LS periodogram peak to be considered real, it must be sufficiently separate in frequency from a significant peak in the periodogram of the window function. Quantitatively, a significant peak in the window function was defined as any peak with power above 20\% of the strongest peak in the window function. The minimum separation in frequency space between a window function peak and an LS peak was chosen to be 0.001 d$^{-1}$. Finally, a false-alarm probability was calculated by performing a bootstrap on the photometric data for the top five peaks that satisfied these conditions. All peaks surviving these quality checks were then considered to be real periodic signals in the data.

To calculate the error in the photometric period corresponding to the best LS periodogram signal for a particular target, a random number of data points were dropped from the dataset and replaced by randomly sampling the original dataset. This new dataset was then analysed as outlined above, and the process was repeated $10^3$ times to generate a distribution of most likely photometric periods, where the final error is quoted as the standard deviation of this distribution. The LS periodograms for the seven targets and their most significant frequency peaks are shown in Figure 2.

### 3.4 Multi-harmonic analysis of photometry

The MHAOV method was then used as an independent check on the LS periodograms. In principle, this method differs from the LS analysis in terms of the model fitted to the data, and the statistic calculated from that fit. Whereas the LS method assumes the data are well modelled by a sine wave, the MHAOV method instead uses orthogonal Szegő polynomials. It should be noted that the first harmonics of the Szegő polynomials simplify to a sine wave, and thus in this work, both the LS and MHAOV method are fitting sine waves to the photometric data.

The advantage of using MHAOV, therefore, comes from how the goodness of fit statistic is calculated (Lachowicz et al. 2006). The statistic used in the LS method is simply $\chi^2$, and depends on the variance of the data. However, because the true data variance is not actually known, but rather estimated from the data, the goodness-

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**Figure 2.** *Left:* Lomb-Scargle periodograms for the six dC stars studied here, in addition to J1250. The most significant peak in the periodogram is highlighted with an orange vertical bar. *Right:* The phase-folded light-curves corresponding to the strongest periodogram peak for each target.
of-fit statistic for the LS periodogram and variance estimations is not independent. The MHAOV method circumvents this problem by using the Fisher analysis-of-variance statistic instead of a $\chi^2$, thus breaking the dependency between the statistic and the variance. This has the benefit of suppressing the significance of periodogram peaks that are due to aliasing in the data; therefore, allowing for the determination of real peaks to be significantly easier than in the LS analysis. Furthermore, the MHAOV analysis is capable of detecting periodic signals at lower S/N than the LS analysis, while also assigning real signals with a higher significance than the LS analysis (Schwarzenberg-Czerny 1996). Based on these statistical advantages, the MHAOV method was used in parallel to the LS method, and in particular to verify the five strongest peaks identified from the LS analysis for each target. The MHAOV analysis agreed with the strongest peak identified through the LS analysis for all targets in this study and hence, only the LS periodograms are shown in Figure 2. All photometric periods quoted in this study are the most significant peak in both analyses.

4 RESULTS FOR INDIVIDUAL STARS

The results of both the photometric and radial-velocity analyses are presented here in detail for each target, with a brief summary provided in Table 3.

4.1 J0842

The dC star J0842 has an unusual spectrum compared to its siblings, as is shown in Figure 3. While the spectrum exhibits clear CN features in the red, below roughly 7000 Å the flux is broadly consistent with a late-K dwarf. However, there are subtle but convincing signs of C$_2$ Swan bands in the blue, especially when compared side by side with an appropriate stellar template of solar abundance (e.g. the late sdK stars SDSS J110140.90+385230.2 or SDSS J124617.53+073749.2). Also noteworthy is the CaH band – a hallmark of metal-poor, cool subdwarfs (Reid & Gizis 2005; Lépine et al. 2007). In addition to H$_\alpha$, close inspection also reveals emission in H$_\beta$ and H$_\gamma$. Without the clear CN features, the spectrum might easily be mistaken for that of a K dwarf, although the emission lines would remain remarkable.

Interestingly, the source is detected with Galex in both the far- and near-ultraviolet bandpasses. There are three and four separate detections respectively, with weighted averages of $m_{\text{NUV}} = 22.0$ AB mag, $m_{\text{FUV}} = 21.4$ AB mag. It is possible that these ultraviolet detections are a result of strong cool-star line emission at these wavelengths, but there are few transitions that could be responsible. In the near-ultraviolet Galex band, the Mg ii 2800 Å doublet is a candidate, but this is highly unlikely. It is near the red cutoff of the bandpass, and furthermore, the emission would have to be three orders of magnitude brighter than the stellar continuum. The corresponding far-ultraviolet detection cannot be due to any common stellar emission features, as the bandpass cuts off before Ly$\alpha$. Therefore, the observed ultraviolet flux is almost certainly not intrinsic to the dC star.

The ultraviolet fluxes appear to be consistent with a white dwarf, and in Figure 3 the ultraviolet and optical photometry for J0842 is approximated using models. Solar-abundance models (Castelli & Kurucz 2003) were used for the dC star, as there are currently no atmospheric models for carbon-enhanced cool dwarfs, while white-dwarf atmospheric models were taken from Koester (2010). A reasonable fit to the ultraviolet-optical photometry is achieved using a $T_{\text{eff}} = 12000$ K, hydrogen-atmosphere white-dwarf model, together with a $T_{\text{eff}} = 4000$ K main-sequence stellar template, where a modestly metal-poor model atmosphere provides a somewhat better fit than one with solar metallicity.

If the white dwarf is at the Gaia parallax distance of $\delta = 680 \pm 50$ pc, then, for the observed temperature and luminosity, evolutionary models suggest a surface gravity close to a canonical value of log $[g (\text{cm s}^{-2})] = 8.0$, with a corresponding mass and radius of 0.61 M$_\odot$ and 0.013 R$_\odot$, and cooling age of around 40 Myr (Fontaine et al. 2001). It is noteworthy that the SDSS u-band photometry appears to be consistent with the composite stellar fluxes, and the Gaia distance indicates clearly that this is a main-sequence star with $M_r = 8.0$ AB mag.

This target was observed during K2 campaign 18. The LS periodogram analysis yields a clearly-defined photometric period of 0.945 $\pm$ 0.001 d with a mean semi-amplitude of $\pm 0.014$ mag (Figure 2).

Nine radial-velocity observations were obtained over three nights in 2019, with two to four measurements per night. Additionally, five further radial-velocity measurements were taken in 2020 Feb. Perhaps surprisingly, these data do not reveal any radial-velocity variations exceeding the error bars (typically $\pm 5$ km s$^{-1}$). Thus, if J0842 is binary with the suspected white-dwarf companion

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Table 3. Summary of period and close binary parameter estimates for dC stars, and possible sources of photometric variability.

| Star    | $P_\text{LS}$ (d) | $K_1$ (Km s$^{-1}$) | $i_{\text{mod}}$ (°) | $\chi^2_{\text{red}}$ | $P_\text{phot}$ (mag) | $\Delta m$ (mag) | Distance (pc) | $M_r$ (AB mag) | $m_{\text{FUV}}$ (AB mag) | Variability? |
|---------|-------------------|---------------------|----------------------|----------------------|----------------------|-----------------|---------------|----------------|---------------------|--------------|
| J0842   | ...               | < 13.8              | < 6.1                | ...                  | 0.95                 | ±0.014          | 680           | 8.0            | 22.0                | +            |
| J0901   | 2.12              | 24.5                | 14.3                 | 2.74                 | 2.12                 | ±0.023          | 590           | 8.2            | ...                 | −            |
| J1015   | 0.20              | 18.5                | 4.9                  | 2.19                 | 0.20                 | ±0.022          | 480           | 8.5            | 18.0                | +            |
| CBS311  | 4.36              | 46.5                | 36.7                 | 0.53                 | 5.26                 | ±0.015          | 100           | 9.2            | 20.3                | −            |
| J1250   | 2.56              | 86.1                | 67.2                 | 0.47                 | 2.92                 | ±0.023          | 280           | 9.2            | > 22.0               | −            |
| SBSS1310| 4.43              | 46.5                | 36.7                 | 0.53                 | 5.26                 | ±0.015          | 100           | 9.2            | 20.3                | −            |
| CBS311  | 0.19              | 29.2                | 7.6                  | 4.83                 | 0.30                 | ±0.030          | 440           | 9.2            | ...                 | +            |

Notes: Errors for the period determinations are non-Gaussian and discussed in detail in Section 4, with posterior distributions shown in the Appendix. The third column lists a model binary inclination based on the adopted or suspected radial-velocity period and velocity semi-amplitude (or upper limit), under the assumptions of 0.4 M$_\odot$ for the dC star, and 0.6 M$_\odot$ for a white-dwarf companion. The fourth column lists the reduced $\chi^2$ for each of the orbital solutions fitted in Figure 1. The final column gives a positive (+) or negative (−) indication as to whether the observed photometric variability amplitude could plausibly result from modulation due to tidal distortion (T), a heating effect from irradiation (I), or rotation (R).

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1 http://www.astro.umontreal.ca/bergeron/CoolingModels
described above, and its orbital period is similar to the observed photometric period of 0.95 d, then the inclination must be ≤ 6.1° (3σ) for masses of $M_{\text{DC}} = 0.4 \, M_\odot$ and $M_{\text{WD}} = 0.6 \, M_\odot$.

### 4.2 J0901

This source exhibits a dC optical spectrum with no evidence for a luminous companion, and is, therefore, a more typical example of its class, albeit with the exception of the H\textalpha and likely H\beta emission features. It is detected in X-rays by Chandra (Green et al., 2019), but Galex observations do not exist for this region of the sky. This study is the first to address its potential binarity, and all radial-velocity observations reported are new.

The J\textsc{oker} radial-velocity analysis indicates that the orbital period distribution is approximately bimodal, with one larger peak at 2.11 d and a second, smaller peak at around 2.09 d. The median of the posterior distribution yields a period of $2.12^{+0.02}_{-0.03}$ d and a velocity semi-amplitude of $24.49^{+1.2}_{-1.0}$ km s$^{-1}$. A modest number of additional radial-velocity measurements at around a phase of 0.7 would likely precisely constrain the orbital period.

The photometric analysis for J0901 was conducted on both ZTF and \textsc{TESS} data. Because of the target faintness, the \textsc{TESS} photometry is not useful, but the ZTF data indicate a possible photometric period of 2.12 d. However, this photometric period remains relatively uncertain.

### 4.3 J1015

This source does not appear in any compendium of dC stars published to date but is flagged as having carbon lines by the SDSS spectroscopic reduction pipeline in \textsc{DataRelease9}. The SDSS spectrum published to date but is flagged as having carbon lines by the SDSS spectrum is the first to address its potential binarity, and all radial-velocity observations reported are new.

The J\textsc{oker} radial-velocity analysis indicates that the orbital period distribution is approximately bimodal, with one larger peak at 2.11 d and a second, smaller peak at around 2.09 d. The median of the posterior distribution yields a period of $2.12^{+0.02}_{-0.03}$ d and a velocity semi-amplitude of $24.49^{+1.2}_{-1.0}$ km s$^{-1}$. A modest number of additional radial-velocity measurements at around a phase of 0.7 would likely precisely constrain the orbital period.

The photometric analysis for J0901 was conducted on both ZTF and \textsc{TESS} data. Because of the target faintness, the \textsc{TESS} photometry is not useful, but the ZTF data indicate a possible photometric period of 2.12 d. However, this photometric period remains relatively uncertain.

### 4.4 CLS 29

The optical spectrum of this dC star is relatively unremarkable, except for the presence of H\textalpha and H\beta emission. The source is possibly marginally detected in the near-ultraviolet with Galex at $m_{\text{NUV}} = 22.9 \pm 0.6$ AB mag, but too uncertain to warrant further analysis. The \textsc{TESS} data yield a clear peak in the periodogram corresponding to a period of 3.62 d. The amplitude of the photometric variation is around 1.4 per cent.

This study is the first to report radial-velocity variations in CLS 29. The radial-velocity analysis is highly constraining, with only 1 in $10^5$ posterior samples surviving. The MCMC analysis yields a bimodal distribution in the orbital parameters, where both peaks in the orbital period are within 1 per cent of 3.14 d. Owing to incomplete phase coverage during the radial-velocity monitoring, the period cannot be uniquely determined. Additional velocity measurements for phases 0.4 – 0.5 should break the degeneracy between the two solutions.

### 4.5 J1250

This dC star is a known single-lined binary and the first shown to have a short-period orbit (Margon et al., 2018), where both photometric and radial-velocity variations are observed. It is included here for context and discussed later, as it belongs to the subgroup of stars exhibiting H\textalpha emission lines, and is detected in X-rays (Green et al., 2019). There are Galex observations for this region of the sky, but no detection near the location of J1250 in either ultraviolet band, placing a rough upper limit at either wavelength of $m > 22.0$ AB mag. For a white-dwarf companion at the \textit{Gaia} parallax distance, this would imply a star no brighter than around $m = 15$ AB mag in the ultraviolet, and for a typical mass and radius this would translate to $T_{\text{eff}} \lesssim 8000$ K.

The publicly available photometric data from ZTF for J1250 were analysed using the LS and MHAOV analyses. The photometric period of 2.92 d identified by Margon et al. (2018) was successfully recovered.

While the radial-velocity measurements of J1250 taken by Margon et al. (2018) do appear consistent with the photometric period, no dedicated radial-velocity analysis was performed in that study. Here, the reported radial-velocity data are analysed using the J\textsc{oker}, where the median posterior distribution for the orbital period is found to be $2.56 ^{+0.11}_{-0.13}$ d, and is notably 12 per cent shorter than the reported photometric period.

### 4.6 SBSS 1310

This is one of the nearest and brightest dC stars known, and among only three or four objects within or just outside 100 pc (Harris et al., 2018). The optical spectrum is spectacular and shows emission from higher Balmer lines as well as C\textsc{ii} H and K (Rossi et al., 2021).

The photometric analysis reveals a strong frequency peak corresponding to a period of 0.196 ± 0.003 d, where two smaller flanking peaks in the power spectrum are the result of windowing. Relatively large photometric errors (averaging 0.07 mag) for this source give rise to the broadened structures in the periodogram. This photometric period is strongest in both the LS and MHAOV analyses, and the phase-folded light-curve at this period shows an amplitude near 2 per cent.
There is a detection of this source in *Galex* in the near-ultraviolet only, with \( m_{NUV} = 20.3 \) AB mag. Similar to the more detailed argument above for J0842, it is unlikely that emission from the dC star is responsible for near-ultraviolet flux detection, and in this case, stellar emission lines would need to be roughly 400 times brighter than the expected continuum. Based solely on the absolute magnitude of this ultraviolet detection, and assuming it was due to a typical surface gravity white dwarf at the Gaia parallax distance to this target, it would imply \( T_{\text{eff}} \approx 7500 \) K.

*Whitehouse et al.* (2018) reported several radial-velocity observations of this star that strongly indicate binarity at a significance just above 7σ. Analysis of the previous and newer radial-velocity data with the *Kepler* yields only 70 surviving samples. The subsequent MCMC posterior distribution for SSBS 1310 is unimodal for all parameters, with the orbital period tightly constrained at 4.43 d with a velocity semi-amplitude of 46.5 ± 0.8 km s\(^{-1}\). The *TESS* data reveal a strong periodic signal corresponding to a photometric period of 5.26 ± 0.01 d. Unexpectedly, the photometric analysis on the individual, cleaned light-curves from each *TESS* sector reveals a clear change in the photometric period; Sectors 15, 16, and 22, individually, yield periods of 5.17 ± 0.01, 5.24 ± 0.01, and 5.30 ± 0.01 d, respectively. Also serendipitously, the amplitude of the photometric variation changes significantly between Sector 16 and 22, by around a factor of five. These data are discussed in more detail in Section 5.2.3.

### 4.7 CBS 311

This system is a previously known binary with a composite spectrum (*Liebert et al.* 1994), consisting of the dC star and a \( T_{\text{eff}} \approx 31,000 \) K DA (= strongest spectral lines are hydrogen) white dwarf. The region of the sky occupied by this target lacks any *Galex* observations that could further constrain the luminosity of the companion. CBS 311 is one of the dC stars detected in X-rays by *Chandra* (*Green et al.* 2019), but it should be noted again that the white dwarf cannot be the source of the high-energy emission. This target was also imaged at high spatial resolution using the *Hubble Space Telescope*, where the binary remained unresolved to at least 25 mas (*Farihi et al.* 2010), corresponding to a projected separation on the sky of roughly 10 AU at the 440 pc *Gaia* DR2 parallax distance.

The dC star in this binary is fainter than suggested by the \( r \)-band brightness due to some flux contribution from the white dwarf. This diminishes the effective S/N of the absorption lines used for cross-correlation when measuring relative and absolute velocities. Therefore, each spectrum of CBS 311 was re-binned to a lower resolution by taking the mean of every 3 pixels. The radial-velocity analysis identifies one prominent peak in the posterior distribution of the orbital parameters, at 0.19 d with small peaks to either side, where the velocity semi-amplitude is constrained to be \( 29.2^{+5.6}_{-4.5} \) km s\(^{-1}\). The photometric analysis was conducted on the available ZTF data for this target. A clear periodic signal corresponding to the strongest periodogram peak is located at 0.302 ± 0.001 d, with a semi-amplitude near ±3 per cent.

### 5 DISCUSSION

Spectroscopic and photometric periods are closely similar for all six systems for which both have been established, and are indistinguishable for only two (Table 3). For those two, orbital variability is, ostensibly, a plausible mechanism for the photometric changes. For the remaining stars (at least), the working hypothesis is that the photometric variability must arise through rotational modulation, with the close similarities between orbital and rotational periods arising through tidal synchronisation.

In this section, binary light-curve simulations are presented, in order to provide context for the observed photometric variability. Potential sources of the periodic flux changes, and their implications for the observed emission-line activity, are then explored. Finally, the origin and evolution of these binaries are discussed in a wider context of carbon-enhanced stars.
Active dC stars with short orbital periods

Figure 4 shows a representative sample of the simulated light-curves (detailed in Section 5.1) for orbital periods representative of the shortest yet measured or suggested for dC stars. The grey bar in each panel represents the range of peak-to-peak variations observed in the dC-star sample (0.028–0.060 mag). The top-left panel demonstrates the influence of tidal distortion alone, while the models represented at top right have identical parameters but include heating (irradiation). The lower-left panel explores the effect of reduced heating (i.e., from cooler white-dwarf companions), and the bottom-right panel shows the effects of decreasing primary mass.

5.1 Binary light-curve simulations

As an aid to interpreting the observed photometric variability, a range of orbital light-curves are simulated in order to gauge the possible effects of close-binary interaction. The goal is to distinguish purely orbital photometric effects from the rotation of the dC star itself. In particular, the aim is to characterise tidal-distortion and heating (irradiation, or day–night) effects arising from a white-dwarf companion, and thereby to characterise the parameter space where these effects may be significant.

A standard Roche model is adopted, with simple ‘deep’ heating, using the current version of a code originally described by Howarth (1982). These illustrative calculations are carried out for the Cousins R band, as a rough match to the observed photometry, with the dC component modelled as a star with $T_{\text{dC}} = 4000$ K, $M_{\text{dC}} = 0.44 M_\odot$, and $R_{\text{dC}} = 0.42 R_\odot$ (if the lobe size allows; lobe-filling otherwise), which are approximately the values expected for a star near the bottom of the K-dwarf sequence for a metallicity [Fe/H] = −1.0 and an age of 1 Gyr or older (Dotter et al. 2008).

The choice of metallicity has a strong influence on the inferred mass and radius for a fixed effective temperature, where going from solar metallicity to [Fe/H] = −2.0 changes the implied stellar mass from roughly 0.6 to 0.3 $M_\odot$ at 4000 K. The metallicity of dC stars remains poorly determined, except in a single case of extreme metal poverty (G77-61 with [Fe/H] = −4.0; Plez & Cohen 2005), yet the kinematics of dC stars clearly indicates that the population contains a large fraction of metal-poor members. The adopted metallicity of [Fe/H] = −1.0 yields $M_r = 8.6$ AB mag, and this compares favourably with the stars studied here, whose absolute magnitudes span $M_r = 8.0–9.2$ AB mag (see Table 3).

The adopted default white-dwarf parameters are $T_{\text{WD}} = 30,000$ K, $M_{\text{WD}} = 0.60 M_\odot$, $R_{\text{WD}} = 0.014 R_\odot$, based on parameters for the white-dwarf companion to CBS 311 (Liebert et al. 1994), one of only two known hot, bright companions to any dC star. The implied mass–radius ratio is fairly typical of single white dwarfs, as well as those found binaries (although in the models the adopted radius is of consequence only for its role in determining the luminosity, $L/L_\odot \approx 0.14$, which determines the magnitude of the irradiation). A white dwarf of this nature gives rise to a significant heating (or day–night) effect on its companion. In order to explore the effect of fainter white dwarfs, each simulation at a given or–period was run for a range of heating efficiencies from 100 per cent down to 20 per cent, where the latter roughly corresponds to $T_{\text{WD}} = 20,000$ K for the same white-dwarf mass and radius.

The models incorporate a number of simplifications that are inconsequential for the purposes of this study. The dC surface intensities are calculated using a constant, linear limb-darkening coefficient of 0.6 (roughly appropriate for the range of temperatures found), a gravity-darkening exponent of 0.08, and emergent fluxes (from Howarth 2011) interpolated at fixed log ($g$ cm s$^{-2}$) = 5.0. While the white dwarf is bolometrically luminous (giving rise to heating of the dC-star atmosphere), it contributes no flux directly to the simulated light-curves. Orbital periods are taken from the range $P_{\text{orb}} = 0.2–1.0$ d – the shortest measured or suggested periods known – in steps 0.1 d, with inclinations $i = 10–90^\circ$ in steps of $10^\circ$.

Figure 4 shows a representative sample of the simulated light-
curves, each panel holding some parameters fixed while varying others. Notably for these parameters, the dC star fills its Roche lobe just below the simulation range at a period of 0.1 d ($R \simeq 0.33R_\odot$). At 0.2 d, the tidal distortion of the star results in a photometric semi-amplitude near ±0.15 mag, and with the characteristic double-peaked, non-sinusoidal pattern. Adding heating at the shortest simulated period results in a single peak and semi-amplitude just over ±0.30 mag. At longer periods (i.e., wider separations), both of these orbital-motion-induced effects diminish rapidly; their full combined effect at 0.4 d is less than half that at 0.2 d. For periods approaching and exceeding 1.0 d, both tidal distortion and irradiation have a negligible impact on the emergent flux.

The magnitude of these calculated effects are taken here only as a representative guide, and they are accurate only to a degree. For the tidal distortion, a classical Roche lobe is used for the equipotential surface, without the influence of asynchronous rotation. The atmosphere of the irradiated star is only changed insofar as a local effective temperature modification, and without any calculation of the actual atmospheric structure. Despite these potential shortcomings, the simulations should be sufficient to determine whether an observed photometric variation can plausibly be attributable to orbital motion, as opposed to stellar rotation.

5.2 Source of photometric variability

Here, the discussion pursues the question of whether the observed photometric variability in any given system can be attributed to orbital motion or to rotation of the dC star.

5.2.1 Rotation versus orbit

The seven dC stars listed in Table 3 exhibit photometric variability amplitudes falling between ±0.014 and ±0.030 mag, with no apparent correlations with orbital period, or with (crudely estimated) inclination. Each system is compared with the light-curve simulations (Fig. 4) for the appropriate orbital period (or range), and none are consistent with ellipsoidal variations induced by tidal distortion. There are two elements to this inference; first, the amplitudes of the predicted changes are significantly smaller than those observed, and secondly, the dominant photometric period of ellipsoidal variations should be exactly half the orbital period. While ellipsoidal variations would occur more readily in systems with shorter orbital periods, neither J1015 nor CBS 311 (each with $P_{\text{orb}} > 0.2$ d) exhibit such light-curve characteristics. As a result of having well-constrained radial-velocity periods for six of the targets, it can be securely stated that none of the light-curves are consistent with an origin in tidal distortions.

Next, the observed photometric changes are compared with those that might be induced due to irradiation from a relatively hot and luminous white-dwarf component – the day–night, or heating, effect. The predicted amplitudes can match or exceed the observed ranges, and so a more detailed examination is needed. The simulations suggest CBS 311 might experience a ±0.15 mag modulation in its light-curve due to irradiation, and the observed value is around five times smaller than this, and thus is unlikely. For J1015, the companion appears to be roughly three times less luminous (see below), and because the irradiated flux scales with luminosity, the heating effect will commensurately decrease. In this case, the simulations indicate that a mildly double-peak structure is expected, with a semi-amplitude near ±0.07 mag, and thus a possibility.

There are two additional effects that would mitigate the magnitude of any orbital-induced photometric changes. First, in the simulations, the white dwarf contributes no flux to the light-curve, and thus in cases where the white dwarf is visible in the optical spectrum, there will be a corresponding dilution factor from the contribution of its flux. In CBS 311, the two stars contribute nearly equal flux in the $r$ band (Liebert et al. 1994), and thus the predicted amplitude of the heating effect on the dC star will be diluted by a factor $\approx 2$. For J1015, the ultraviolet through optical spectra and photometry are decently fit by the addition of a $T_{\text{eff}} = 23500$ K DA-type white dwarf, whose flux in the $r$ band is approximately 25% of the total, and the amplitude of any heating effect will be reduced accordingly. Second, for an $i = 0^\circ$ (face-on) binary, there is no expected flux modulation from either tidal distortion or irradiation, as both binary components are in full view for the entire orbit. The amplitude of both these orbital effects on the light-curve should increase to first order as $\sin i$ (Figure 4), and thus in real observations, where present, these effects will typically be modestly decreased.

Taking into account all of these factors as the data permit, the light-curves of both CBS 311 and J1015 become inconsistent with irradiation for relatively low inclinations. Specifically, for $i < 42^\circ$ the observed photometric variation of CBS 311 is more consistent with rotation, and similarly for J1015 for $i < 25^\circ$. For the stars J0842, J0901, CLS 29, J1250, and SBSS 1310, the only realistic source of photometric variability is rotation. This is next considered as the primary or exclusive cause of the observed light-curve variations in the sample of dC stars is axial rotation.

5.2.2 Tidal synchronisation

Figure 5 shows the relationship between the photometric and orbital periods for the six targets with well-constrained radial-velocity periods (including J1250). Two systems have essentially identical photometric and radial-velocity periods, while the remaining four have photometric periods that are modestly longer than the radial-velocity periods (further undermining the possibility that the photometric variations are orbital).

Those systems with matching photometric and radial-velocity
periods – J0901 and J1015 – are neither the shortest period systems, nor those with the most luminous companions, as might be expected if tidal distortion or irradiation were responsible for the photometric variability. The stars J1250, CLS 29, SBSS 1310, and CBS 311 possess photometric periods on average around 15 per cent longer than their orbital periods, which is strictly inconsistent with orbital motion.

The collective data on these emission-line dC stars therefore broadly support a rotational origin for the photometric changes. However, magnetic braking causes spin-down with age for isolated convective stars. Assuming a crude, and probably conservative, age estimate for dC stars of a few to several Gyr (Farhi et al. 2018), the empirical age–rotation relationship for single M dwarfs predicts spin periods on the order of 20–100 d (Engle & Guinan 2018) – significantly longer than any of the photometric periods measured in this study, implying an additional source of spin angular momentum.

Tidally induced rotational synchronisation is an appealing mechanism for the implied spin-up and relatively rapid rotation, particularly since the radial-velocity data are consistent with circular orbits (and rotational synchronisation is expected to take place on shorter timescales than orbital circularisation). Theoretical timescales for synchronisation depend mainly on the ratio of the radius of the star experiencing tidal forces (here the dC star) to the orbital separation between the binary components. Adopting canonical values for M-dwarf radii, along with the orbital separations determined from this study, it is found that synchronisation timescales are a few Myr up to no more than a Gyr (Zahn 1989, 2008).

Assuming that all of the binary companions are now white dwarfs, their cooling timescales can be estimated and provide a window over which tidal synchronisation could have occurred. Depending on the Gaia ultraviolet flux in these systems to white dwarfs – and upper limits for non-detections – at the Gaia distances to the dC stars, and assuming a fixed white-dwarf surface gravity of log \( g \) (cm s\(^{-2}\)) \( = 8.0 \), cooling-age estimates and lower limits can be derived from models (Bédard et al. 2020). The resulting cooling ages range from approximately 10 Myr (CBS 311) to over 1 Gyr, but in all cases are broadly consistent with the predicted synchronisation timescale.

### 5.2.3 Differential rotation

Of course, for fully synchronised rotation the photometric signal should exactly match the radial-velocity period (as assumed in the analysis of J1250 by Margon et al. 2018). However, there are modest offsets between the orbital and photometric periods for much of the sample, with the photometric period never shorter than the radial-velocity period. To reconcile this result with fully synchronised equatorial rotation, the dC star must rotate differentially, with starspots (or other surface brightness inhomogeneities) at relatively high latitudes; for example, solar differential rotation implies a 15 per cent longer period at latitudes of 20–35° than at the equator.

The plausibility of differential rotation is best illustrated by SBSS 1310. As noted in Section 4.6, the photometric period of this star shows modest but significant changes between TESS Sects 15, 16, and 22 (5.17, 5.24, and 5.30 d, respectively, with \( \sigma \approx 0.01 \) d uncertainties). The individual sector periodograms are presented in Figure 6, together with the sector light-curves phase-folded to the mean 5.26 d period. Evidently, the light-curve changes not only in period, but also in shape and amplitude. The simplest interpretation of these characteristics is changes in the geometry of starspots (resulting in the variable light-curve shape and amplitude) and in their latitude, accompanied by differential rotation (resulting in the variable period). Differential rotation also accommodates the difference between the radial-velocity and mean photometric periods (4.4 d vs. 5.26 d), suggesting a star-spot latitude of \( \gtrsim 20° \) (assuming solar-type differential rotation and equatorial synchronisation).

If SBSS 1310 is representative of its spectral sub-class, then the targets appear to be differentially rotating, tidally synchronised stars.

### 5.3 Source of stellar activity

There are two basic possibilities for the source of stellar activity in dC stars, as manifested by H\(\alpha \) and similar emission features: an intrinsic source via relatively rapid stellar rotation, and an extrinsic source via irradiation.

In order to heat a low-mass star to sufficiently high local surface temperatures that emission lines are excited, a white-dwarf companion must be close to the irradiated star and fairly luminous. Based on a comprehensive study of stellar activity as traced by H\(\alpha \) emission in M dwarf companions to white dwarfs (Rebassa-Mansergas et al. 2013), it was found that activity is nearly always intrinsic for both partly and fully convective stars. Furthermore, while the influence of magnetic braking is seen for partly convective stars in wide binaries, those M dwarfs in short-period orbits resulting from a common envelope all remain active. This fact alone argues that irradiation is not the cause of stellar activity in low-mass stars closely orbiting white dwarfs, and the study furthermore demonstrates that only white dwarfs with an irradiating flux that is at least 3× higher than the intrinsic M dwarf flux can result in emission (Rebassa-Mansergas et al. 2013). None of the seven short-period dC stars in this study approaches this level of irradiation; e.g. the 30 000 K white dwarf in CBS 311 is only 1.5× more luminous than its companion, and thus falls drastically short of this criterion at the surface of the dC star. Therefore, irradiation cannot be responsible for the observed activity in dC stars.

The remaining possibility requires relatively rapid rotation to generate a solar-type dynamo (Parker 1955), and an active chromosphere (Boro Saikia et al. 2018). The fact that dC stars with emission lines are relatively rapid rotators is clearly evidenced by X-ray detections in six out of six observed cases (Green et al. 2019), including four of the stars analysed here. There are three potential angular momentum sources for the rapid rotation observed in dC stars: intrinsic spin of formation, spin-up owing to mass transfer, and tidal synchronisation in a binary.

Three lines of evidence discount the idea that dC stars are relatively young and hence, still rotating rapidly due to their youth. One is their overall kinematical properties that suggest a mixture of old disc and halo stars (Farhi et al. 2018), second is the indication from population synthesis that metal-poor stars are more readily polluted to C/O \( > 1 \) than solar abundance stars (de Kool & Green 1995), and third is through estimating the cooling ages of their white-dwarf companions. The actual masses and temperatures of dC stars are certainly not well constrained, and not established yet at any level of confidence, but based on their positions on the H–R Diagram they should be partly convective stars with parameters comparable to late-type dwarfs, and because of inevitable magnetic braking (Skumanich 1972; McQuillan et al. 2013), they require an extrinsic origin to their current spins, assuming their total ages are at least a few Gyr.

Thus, only mass transfer or tidal synchronisation remain as possibilities for the observed, spin-induced stellar activity. And it
should be noted that both can occur, that is they are not mutually exclusive (however, it is noteworthy that the latter can erase any evidence of the former, e.g. in the case the dC was once spinning faster than the orbital period). If some fraction of the emission-line dC stars are the product of spin-up through mass transfer, and not the result of tidal synchronisation in short-period binaries, the observed rotation rate would be an indication of how much mass the system has accreted. Therefore, measurements of the rotation rates in any systems that are not short-period binaries (such as for the Ba giant HD 165141; Theuns et al. 1996) would provide constraints on the amount of mass, and angular momentum accreted, and thus inform mass transfer models (Izzard et al. 2010; Matrozis et al. 2017).

5.4 Period distribution

The overall results indicate clearly that Hα emission in dC stars is linked to short-period binarity. In fact, of the ten dC stars with constrained orbital parameters, there are two distinct families: those with Hα emission all have orbital periods of a few days or less, those without detectable Hα emission have orbital periods of several months or longer (Dearborn et al. 1986; Harris et al. 2018; Margon et al. 2018).

The data at hand suggest that the orbital period distribution of the dC star population may be bimodal, and thus consistent in some key aspects with the (vastly) more commonly occurring M-type dwarf companions to white dwarfs. In general, such bimodal populations are consistent with the post-main sequence removal of intermediate-period orbits, via a combination of common-envelope evolution for shorter periods, and the reduction of gravitational binding energy due to mass loss for longer periods. However, there is a notable difference in the dC star binary population, and that is the orbital periods on the order of years, which are absent in the carbon-normal, M-dwarf companion population orbiting white dwarfs (Farhi et al. 2010; Nebot Gómez-Morín et al. 2011; Ashley et al. 2019). It is tempting to speculate that the dC companion orbits may differ from dM companion orbits owing to mass transfer, but that is beyond the scope of the current study. Nevertheless it is clear that the short-period dC binaries have experienced a common envelope (Izzard et al. 2012; Ivanova et al. 2013). If the emission-line dC stars have evolved through a common-envelope phase, and assuming their carbon enhancement is extrinsic, then this suggests relatively stable mass transfer prior to the common envelope.

Extrinsically carbon-enhanced binaries possessing short-orbital periods are rare. Amongst the Ba, CH, and CEMP-s populations, which typically possess orbital periods of many months or years (Lucatello et al. 2005; Jorissen et al. 2016, 2019), only two stars are known to have orbital periods shorter than ten days. One of these is the central binary in the Necklace nebula, with a 1.2 d orbital period, and whose exact link to other carbon-enhanced populations remains to be determined (Miszalski et al. 2013). The other is HE 0024-2523, a 0.9 M⊙ main-sequence star with a 3.4 d orbital period measured through radial-velocity monitoring (Lucatello et al. 2005), and a CH spectral class. Similar to the short-period binaries found here, HE 0024-2523 is suspected of receiving mass transfer from an evolved companion before evolving through a common envelope. However, there is no evidence of short-period photometric variability or rotation in HE 0024-2523, despite decent quality TESS data, and the star lacks emission lines.

It is tempting to speculate why, on the one hand, Ba, CH, and CEMP-s stars appear to result from relatively wide and essentially detached binary evolution; and on the other hand, why the vast majority of post-common envelope companions to white dwarfs appear to be carbon-normal M dwarfs. Unfortunately, there are currently no robust constraints on the fundamental stellar parameters of dC stars (i.e. mass, effective temperature, metallicity), and a distinct lack of appropriate atmospheric models. It is therefore beyond the scope of the current study to address the different pathways in stellar and binary evolution that might give rise to the observed orbital period differences in these populations. However, it is hoped that the empirical results presented here will encourage novel evolutionary modeling. Because dC stars should far outnumber their evolved counterparts, those with short orbital periods provide valuable constraints for models of extrinsic carbon-enrichment.

6 CONCLUSIONS

This study analyses time-series radial-velocity and photometric data for seven dC stars with Hα emission lines. Six of the dC stars
analysed here are found to possess orbital periods in the range 0.2–4.4 d, along with photometric periods that are either identical (in the case of J0901 and J1015), or within a few tens of per cent of the orbital period. The seventh dC star, J0842, lacks significant radial-velocity variations but possesses a 0.945 d photometric period, and if these periods are similar it would be difficult to detect through ground-based radial-velocity measurements.

Binary light-curve simulations suggest that the origin of the photometric variations observed in these emission-line dC stars is consistent with rotation. None of the light-curves show any evidence for a day- and night-side effect or tidal distortion, and while irradiation is physically plausible in two of the seven systems, it is unlikely based on binary inclination and companion brightness estimates. Thus, rotation is the most likely cause for the observed photometric variations. The mostly sinusoidal behavior of the (best) observed light curves is not necessarily expected for a random assortment of starspot latitudes and stellar spin inclinations, but both the light-curve data quality and the number of short-spin period dC stars are insufficient to determine further constraints on the geometry and number of starspots.

The expectations of magnetic braking are such that the dC stars should be spinning an order of magnitude more slowly than observed, and their ages are consistent with sufficient time to have become tidally synchronised to their visible or unseen (and presumed to be) white-dwarf companions. Tidal synchronisation provides the most consistent picture for these short-period dC star binaries, and provides the additional angular momentum needed for spin-up, which in turn powers the emission lines through a solar-type dynamo. However, mass transfer cannot be ruled out, but in the sample studied here, mass transfer is only necessary for the extrinsic source of pollution, and not for spin-up. Lastly, at least a few of the dC stars appear to be differentially rotating, as evidenced by photometric periods that are similar to, but somewhat longer than their orbital periods, as well as at least one system with changing photometric periods that are similar to, but somewhat longer than their orbital periods, as well as at least one system with changing photometric period and variability amplitude.

This study more than doubles the number of dC stars with constrained orbital periods, and increases those known to be at short-period orbits from one to six. It appears the dC binary period distribution may be bimodal, with a post-common-envelope population as studied here, and a wider population with periods on the order of months to years. Interestingly, the other well-studied, carbon-enhanced stellar populations lack a short-period, post-common-envelope population, and it is likely the dC stars will provide key insight into mass transfer models on the basis of these, and future orbital period determinations.

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Software: Astropy (Astropy Collaboration et al. 2018), Corner (Foreman-Mackey 2016), iraf (Tody 1993), Joker (Price-Whelan et al. 2020), Matplotlib (Hunter 2007), NumPy (Harris et al. 2020), Pandas (McKinney et al. 2010), SciPy (Virtanen et al. 2020)

DATA AVAILABILITY STATEMENT

TESS and K2 data used in this article are available via the Mikulski Archive for Space Telescopes portal, maintained and funded by NASA. ZTF data are available to download in the NASA/IPAC infrared science archive. Finally, ISIS spectroscopy are available at the Isaac Newton Group Archive which is maintained as part of the CASU Astronomical Data Centre at the Institute of Astronomy, Cambridge.

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APPENDIX A: POSTERIOR DISTRIBUTIONS FOR RADIAL VELOCITY DETERMINATIONS

This paper has been typeset from a LaTeX file prepared by the author.
Figure A1. The radial-velocity posterior PDFs for all targets analysed using the joker rejection sampler presented in order of right ascension; J0842, J0901, J1015, CLS 29, J1250, SBSS 1310, and CBS 311. The orbital period, velocity semi-amplitude, and systemic velocities are presented along with the 16 per cent and 84 per cent confidence limits for each target.
$P \ [d] = 2.56^{+0.13}_{-0.11}$

$J1250$

$K \ [\text{km s}^{-1}] = 86.17^{+4.21}_{-3.42}$

$P \ [d] = 4.43^{+0.00}_{-0.00}$

$SBSS\ 1310$

$K \ [\text{km s}^{-1}] = 46.52^{+0.84}_{-0.84}$

$P \ [d] = 0.19^{+0.00}_{-0.00}$

$CBS\ 311$

$K \ [\text{km s}^{-1}] = 29.19^{+2.03}_{-1.73}$

Figure A1 – continued