Connecting low- and high-mass star formation: the intermediate-mass protostar IRAS 05373+2349 VLA 2

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ABSTRACT

Until recently, there have been few studies of the protostellar evolution of intermediate-mass (IM) stars, which may bridge the low- and high-mass regimes. This paper aims to investigate whether the properties of an IM protostar within the IRAS 05373+2349 embedded cluster are similar to that of low- and/or high-mass protostars. We carried out Very Large Array as well as Combined Array for Research in Millimeter Astronomy continuum and \(^{12}\) CO(J=1–0) observations, which uncover seven radio continuum sources (VLA 1–7). The spectral index of VLA 2, associated with the IM protostar is consistent with an ionized stellar wind or jet. The source VLA 3 is coincident with previously observed H\(_2\) emission line objects aligned in the north–south direction (P.A. \(-20\) to \(-12^\circ\)), which may be either an ionized jet emanating from VLA 2 or (shock-)ionized cavity walls in the large-scale outflow from VLA 2. The position angle between VLA 2 and 3 is slightly misaligned with the large-scale outflow we map at \(~5\)-arcsec resolution in \(^{12}\) CO (P.A. \(\sim30^\circ\)), which in the case of a jet suggests precession. The emission from the mm core associated with VLA 2 is also detected; we estimate its mass to be 12–23 M\(_{\odot}\), depending on the contribution from ionized gas. Furthermore, the large-scale outflow has properties intermediate between outflows from low- and high-mass young stars. Therefore, we conclude that the IM protostar within IRAS 05373+2349 is phenomenologically as well as quantitatively intermediate between the low- and high-mass domains.

Key words: techniques: interferometric – stars: formation – stars: protostars – ISM: jets and outflows – radio continuum: stars.

1 INTRODUCTION

As much research has focused on either distinguishing or unifying the formation of high-mass stars and their low-mass counterparts, intermediate-mass (IM) stars are of interest as they provide a bridge between the two mass regimes. IM protostars – the precursors of Herbig Ae/Be stars – are defined as young stellar objects (YSOs) that will reach final masses of 2–8 M\(_{\odot}\) and have luminosities between \(~50\) and 2000 L\(_{\odot}\) (Beltrán 2015). The lower limit of 2 M\(_{\odot}\) originates from the fact that above this limit stars should not have the outer convective zones required to produce the magnetic fields needed for magnetically-mediated accretion (see Simon et al. 2002, for an observational confirmation of this temperature and thus mass limit); the upper limit of 8 M\(_{\odot}\) originates from the fact that above this limit stars should not have the outer convective zones required to produce the magnetic fields needed for magnetically-mediated accretion (see Simon et al. 2002, for an observational confirmation of this temperature and thus mass limit); the upper limit of 8 M\(_{\odot}\) originates from the stellar mass required to produce a type II supernova, which is used to define the lower mass limit for high-mass stars (Zinnecker & Yorke 2007). The upper mass limit also corresponds to the mass above which photoionization by UV photons becomes easily observable in cm continuum and the mass above which a pre-main-sequence phase is not observable (Beltrán & de Wit 2016).

Unlike high-mass stars, whose descent on to the main sequence occurs while still deeply embedded and actively accreting within their parent cores, IM stars have a longer pre-main-sequence time-scale, so that their discs are revealed for part of their formation, similar to their low-mass counterparts. This is a result of their accretion time-scales being shorter than their Kelvin–Helmholtz time-scales, so that these stars are finished accreting before they reach the main sequence (Beuther et al. 2007). On the other hand, IM stars form in more densely clustered environments, similar to high-mass stars (e.g. Gutermuth et al. 2005; Fuente et al. 2007), with a smooth transition to the regime of rich clusters at a mass of 6 M\(_{\odot}\) (Testi, Palla & Natta 1999). Thus, IM protostars should also share some of the characteristics of those at higher masses.

To date, there have been only a handful of high-resolution studies of IM protostars. One of the first, Fuente et al. (2001) presented a study including two IM protostars: NGC 7129 FIR1 and FIR2, uncovering multiplicity and/or associated clusters, protostellar envelope masses of 2–3.5 M\(_{\odot}\) and energetic bipolar outflows which appear to be driven by several sources. Beltrán, Girart & Estalella

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(2006) studied the morphology of the outflow from the IM protostar IRAS 22272+6358A, finding that the 14.2 M_☉ core OVRO 2, one of the four detected continuum sources in the region, was powering the asymmetric and collimated molecular outflow. The properties of the outflow were consistent with those of the outflows driven by low-mass YSOs. In a companion paper to Beltrán et al. (2002, 2006) that studied IRAS 21391+5802 and IRAS 22272+6358A mentioned above, Beltrán et al. (2008) studied the 280 L_☉ protostar IRAS 20050+2720. They detected three dust cores in the region at 2.7 mm (OVRO 1-3), and two bipolar flows in 12CO J = 1 – 0, one of which is powered by OVRO 1. Combining their study of the outflows from these three IM protostars with a small number of objects from the literature, they concluded that although outflows from IM protostars are more energetic than those from low-mass protostars, they were not necessarily more complex, being collimated even at low velocities. In addition, they suggest that the increased outflow momentum rates compared to outflows from low-mass protostars are likely due to higher accretion rates for these objects. A further finding was that the IM protostars studied all formed in protoclusters containing several IM and low-mass millimetre sources.

More recently, van Kempen et al. (2016) used APEX to study a sample of six IM protostars, mapping the bipolar outflows and quiescent gas in 12CO and 13CO. They showed that their line luminosities and outflow forces also follow trends with bolometric luminosity and outflow mass, connecting low- and high-mass protostars, however they could not confirm the result of Beltrán et al. (2008) that fragmentation enhances outflow forces for IM protostars in clusters. Studying one of the largest samples thus far, Crimier et al. (2010) determined the physical structure of the envelopes of a sample of five IM protostars via radiative-transfer modelling, finding that the physical parameters describing the envelopes vary smoothly between low- and high-mass protostars, and that the density structure was consistent with predictions from the predictions of ‘inside–out’ collapse.

Other IM protostars studied to date include IRAS 22198+6336 (Sánchez-Monge et al. 2010; Palau et al. 2011), AFGL 5142 (Palau et al. 2011), IC 1396 N (Neri et al. 2007; Fuente et al. 2009), 13 S in R CrA (Saum 2015), MMS 6/OMC-3 (Takahashi & Ho 2012; Takahashi et al. 2012), IRAS 05345+3157 (Fontani et al. 2009) and G173.58+2.45 (Shepherd & Watson 2002). In this work, our aim is to increase this relatively small number by investigating the physical properties of the IM protostar associated with IRAS 05373+2349, which has a luminosity of 290–600 L_☉ (Molinari et al. 2000; Zhang et al. 2005; Kanzhazydan et al. 2011), to uncover the differences and similarities to low- and high-mass star formation. To do this, we have carried out the first mm-interferometric observations of IRAS 05373+2349, which lies at a distance of ~1.2 kpc (Molinari et al. 2000). The associated embedded cluster and the nearby GGD 4 object have been the focus of several studies due to the active star formation occurring within the region. Zhang et al. (2005) has mapped the large-scale outflow from IRAS 05373+2349 in 12CO (J=1–0) with 29-arcsec resolution, determining that it has an NE–SW orientation. A more recent study by Kanzhazydan et al. (2011, hereafter K11) identified numerous possible outflows in the immediate area of IRAS 05373+2349 based on near-infrared H2 line emission and investigated the candidate driving sources.

In this paper, we present continuum observations of IRAS 05373+2349 at 6, 3.6, 1.3-mm and 2.7-mm wavelengths as well as 12CO (J=1–0) line emission. In Section 2 we describe the observations. In Section 3 we present our results, including the derived outflow properties. We present our discussion in Section 4 and in Section 5 we outline our conclusions.

## 2 Interferometric Observations

### 2.1 VLA 6, 3.6 and 1.3-cm continuum

Multi-wavelength radio continuum observations were taken of IRAS 05373+2349 over two days during 2008 March, with the Very Large Array (VLA) of the National Radio Astronomy Observatory.¹ The program number for these observations was AJ337. The pointing centre was 05°40′24.40 +23° 50′54.00 (J2000).

The first set of observations in the $\lambda (\lambda = 3.6$-cm) and $K (\lambda = 1.3$-cm) bands were taken on 2008 March 10 and the observation of the C($\lambda = 6$-cm) band was taken on 2008 March 12. For each observation two 50-MHz spectral windows were placed at: 4.89 and 4.84 GHz, 8.44 and 8.49 GHz, and 22.5 and 22.4 GHz for the 6, 3.6 and 1.3-cm bands, respectively. The observations were performed in the VLA’s C configuration which had baseline lengths between 35 m and 3.2 km, which produced information on angular scales from 3.88, 2.31 and 1.4 arcsec to 4.9, 3.0 and 1.4 arcmin for 6, 3.6 and 1.3 cm, respectively.

Table 1 presents a summary of the observations which lists the observed wavelength, configuration, observation date, number of antennas (with the number of antennas with useful data given in parentheses), time on-source, synthesized beam size, position angle (P.A.) and the map rms noise.

Table 2 presents the calibrators used in the observations; the ‘Cal type’ column showing their application. Table 2 also presents the fluxes derived from the available models for the flux calibrators 3C48 (0137+331) and 3C286 (1331+305), and the bootstrapped fluxes derived from the gain calibrators.

Data reduction and imaging was carried out using the Common Astronomy Software Applications (CASA) package version 4.2.2 (McMullin et al. 2007). Briggs weighting with a robust parameter of 1.5 was used to clean the images, to retain optimal sensitivity in the observations. Data from baselines shorter than 50 m were removed from the 6-cm image to eliminate flux that was partially resolved-out.

### 2.2 CARMA 12CO (J=1–0) and 2.7-mm continuum

Combined Array for Research in Millimeter Astronomy (CARMA²) observations at $\lambda = 2.7$ mm were taken on 2007 April 24. The antenna array during the observation consisted of 15 antennas, six 10 m and nine 6 m in diameter with primary beams sizes of 64 arcsec and 115 arcsec, respectively. A mosaic was created from nine pointing centres in a square grid pattern. The pointings were separated by 30 arcsec.

The CARMA correlator was set up with two sidebands, placed either side of the chosen local oscillator frequency of 113.280 GHz, with the upper bands situated to measure 12CO (J=1–0) line emission. Both the upper and lower sidebands contained one wide and two narrow spectral windows, giving a total of six windows. The

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wide spectral windows each had bandwidths of 468.750 MHz, a
total of 15 channels with a spectral resolution of 31.250 MHz and
were centred on 111.558 and 115.000 GHz, respectively. The two
narrow windows in each sideband had bandwidths of 30.76 and
7.69 MHz, 63 channels and spectral resolutions of 488.281 and
122.070 kHz. The narrow windows were centred on 111.288 GHz
in the lower sideband and 115.271 GHz in the upper sideband.

Hamming smoothing and line length corrections were applied to
the data in MIRIAD, which gave spectral resolutions of 1.27 and
0.32 km s\(^{-1}\) for the two narrow windows covering the \(^{12}\)CO
(J=1–0) line emission. The data were then exported from their
original format into CASA where data reduction and imaging were
conducted. The source 0530+135 was used to calibrate both ampli-
tude and gain of the observation; the assumed flux is listed in
Table 2.

The images created from the millimetre continuum and \(^{12}\)CO
(J=1–0) line emission are presented in Section 3.2. The two con-
tinuum spectral windows were cleaned together to produce a sin-
gle image with a central frequency of 113.279 GHz, using Briggs
weighting with a robust parameter of 0.5. The synthesized beam
size and noise, along with the number of antennas and time spent
on-source for the observation, can be found in Table 1.

For the \(^{12}\)CO (J=1–0) line image, the two narrow spectral
windows were combined at 115.271 GHz using natural weight-
ing and a channel width of 1.3 km s\(^{-1}\); uv–distances smaller than
3 k\(^{\lambda}\) were excluded. The beam size of the \(^{12}\)CO (J=1–0) image
was 4.92 \(\times\) 4.28 arcsec P.A. 78\(^\circ\) and the noise ranged
from \(\sim\)0.07 Jy beam\(^{-1}\) in an empty channel, up to \(\sim\)0.5 Jy beam\(^{-1}\)
in the inner channels.

### 3 RESULTS

**3.1 VLA 6, 3, and 1.3-cm continuum**

In this section, we present the VLA images at each wavelength and
calculate spectral indices.

The radio continuum maps are presented in Fig. 1, which shows
the 6 and 3.6-cm continuum emission with a wide view of the
surrounding region (panels a and b), and towards VLA 2 and 3
(panels c and d), and in Fig. 2, which shows the 1.3-cm continuum
towards VLA 2.

Table 3 presents the peak position, integrated and peak flux den-
sities, the deconvolved source size and position angle and the con-
volved size measured as the largest extent of the 3\(\sigma\) contours. The
properties presented in Table 3 were derived from fitting a 2D Gauss-
ian to each radio source, with the exception of two non-Gaussian
sources which are indicated by an * in the \(\lambda\) column.

For Gaussian sources unresolved by the fit, an upper size limit
is given in the deconvolved source size column of Table 3. The
upper size limit is determined by calculating the maximum possible
source size when unresolved, which depends on the signal-to-noise
of the source and the synthesized beam size of the observation.

For non-Gaussian sources, the peak flux density and position
were measured at the brightest pixel and the integrated flux density
was measured by summing up the emission within a region defined
by the 1\(\sigma\) contour of the source.

The fitted source positions had an average error of \(\sim\)0.1 arcsec.
The errors in the integrated flux of the non-Gaussian sources were
determined by measuring the noise in nearby empty regions of sky
using the same aperture shape used for the flux measurements.

The spectral index, \(\alpha\), defined as

\[
S_\nu \propto \nu^\alpha
\]

(1)

was determined for all sources where there were fluxes at more
than one wavelength. In the case of VLA 2, which had three fluxes,
the spectral index was derived from a least squares fit in log space,
including errors in the fit. The spectral indices are given in Sec-
tions 3.1.1 and 3.1.2 below. The angular size–frequency index, \(\xi\),
given by

\[
\theta_\nu \propto \nu^\xi
\]

(2)

The two indices (\(\alpha\) and \(\xi\)) provide further classification of the
emission coming from a thermal or non-thermal source. For ther-
mal sources with an \(\alpha = 0.6\), which is related to the power-law
relationship between electron density and emitting object radius,
it can be shown that \(\xi = -0.7\) (Panagia & Felli 1975). Whereas,
non-thermal synchrotron emission, detected using interferometric
techniques, can produce \(\xi \leq -2\) (Thompson, Moran & Swenson
1986). We determined \(\xi\) for VLA 1, which was resolved at more
than one wavelength, using a similar fitting method as described
above.

#### 3.1.1 VLA 2 and VLA 3

VLA 2 is associated with one of the most prominent members of the
embedded cluster towards IRAS 05373+2349 and was detected at
all wavelengths. Figs 1(c) and (d) shows the 6 and 3.6-cm emission,
respectively. VLA 2 was the only source detected at 1.3 cm, which
is shown in Fig. 2.

The markers in these figures show: the peak position of the mil-
limetre core associated with VLA 2 and 3 detected with CARMA

| Wavelength | Array | Date of observation | No. of antennas | Time on-source (h) | Synthesized beam size (arcsec) | P.A. (°) | Map rms (mJy beam\(^{-1}\)) |
|------------|-------|---------------------|----------------|------------------|-------------------------------|---------|-----------------------------|
| 6 cm       | VLA-C | 2008 Mar 12         | 27 (24)        | 1.5              | 4.55 \(\times\) 4.02          | 27      | 0.026                       |
| 3.6 cm     | VLA-C | 2008 Mar 10        | 26 (23)        | 2.0              | 2.97 \(\times\) 2.86          | –13     | 0.016                       |
| 1.3 cm     | VLA-C | 2008 Mar 10        | 26 (21)        | 1.3              | 1.23 \(\times\) 1.04          | 48      | 0.047                       |
| 2.7 mm     | CARMA-D | 2007 Apr 24    | 15             | 5.7              | 4.92 \(\times\) 4.28          | 78      | 2.6                         |

### Table 1. Summary of continuum observations.

| Source          | 6 cm | 3.6 cm | 1.3 cm | 2.7 mm | Cal type |
|-----------------|------|--------|--------|--------|----------|
| 0530+331        | 5.52 | –      | –      | –      | A        |
| 1331+305        | –    | 5.22   | 2.58   | –      | A        |
| 0530+135        | –    | –      | –      | 6.70   | A+G      |
| 0559+238        | 0.41 | 0.37   | –      | –      | G        |
| 0539+145        | –    | –      | 0.35   | –      | G        |

**Note.** Absolute flux calibrators are denoted by ‘A’ and the gain calibrators
are denoted by ‘G’ in the ‘cal type’ column.

wavenumbers - index - predicted - calculated - deviations - uncertainties.
Figure 1. Continuum emission at 6 cm (left column, $\sigma = 26 \, \mu$Jy beam$^{-1}$) and 3.6 cm (right column, $\sigma = 16 \, \mu$Jy beam$^{-1}$). Panels a and b show the region surrounding the IRAS source. Panels c and d show VLA 2 and 3. The position of the millimetre core seen at 2.7 mm (see Fig. 3) associated with VLA 2 is marked by a plus. A Class I object reported by Gutermuth et al. (2009) is marked by a black star. The black xs indicate the positions of MHO 745 A, B and C detected by K11. Contour levels show for (a and b): $-3, 3, 4, 5, 6, 8, 10, 12, 15 \times \sigma$, (c): $-3, 3, 4, 5, 6, 8, 10, 12, 15 \times \sigma$, (d): $-3, 3, 4, 5, 6, 8, 10, 12, 15, 20, 25 \times \sigma$. The beam and scale bar are shown at the bottom of the images.

(see Section 3.2), marked as a plus sign; a Class I protostar (to the south–east of VLA 2, Gutermuth et al. 2009), marked as a black star, and the positions of three molecular hydrogen emission line objects (MHO 745A, B and C, K11), marked as crosses.

The spectral index of VLA 2 was $\alpha = 0.38 \pm 0.14$, measured from fitting a line to the three fluxes$^3$ given in Table 3. As the 1.3-cm image, shown in Fig. 2, is sensitive to more compact emission than 6 and 3.6-cm maps, while being less sensitive to the extended emission, it may be that we do not detect a significant portion of the flux of VLA 2 at this wavelength. Therefore, it is possible that the spectrum including the 1.3-cm flux measurement is artificially flattened. Ignoring the 1.3-cm flux yields a much larger spectral index of $\alpha = 1.20 \pm 0.24$. Both of these spectral indices are consistent with free–free emission of varying degrees of optical thickness.

VLA 2 was detected previously at 3.6 cm with a flux density of $0.70 (\pm 0.04) \, \text{mJy}$ and an estimated spectral index of 0.9 (Molinari et al. 2002). Our measured flux and spectral index for VLA 2 is consistent with this finding.

VLA 3 is a new radio source, 6.7 arcsec and 7.9 arcsec (at 6 and 3.6 cm, respectively) to the north of VLA 2. These correspond to separations of $\sim 8000$ and $\sim 9500 \, \text{au}$ at 1.2 kpc. VLA 3 was resolved in our 3.6-cm observation, being extended roughly north–south with a P.A. of 15°. The spectral index for VLA 3 was found to be $\alpha = 0.45 \pm 0.39$ indicating partially optically thick free–free emission. We find no UKIRT Infrared Deep Sky Survey (UKIDSS; Lawrence et al. 2007) emission associated with this source, which would indicate a separate protostar.

$^3$The integrated and peak flux densities of VLA 2 at 6 cm, measured in a 1$\sigma$ region, was $0.43\pm0.06 \, \text{mJy}$ and $0.41\pm0.03 \, \text{mJy beam}^{-1}$, respectively. These are very similar to the fluxes derived from the Gaussian fit given in Table 3, thus the determined spectral index is not affected by different methods of flux measurement.
Figure 2. A close-up map of the 1.3-cm continuum emission towards VLA 2 with the position of the millimetre core seen at 2.7 mm (see Fig. 3) associated with VLA 2 marked by a plus. Contour levels are $-3, 4, 5, 6, 7, 8$ and $9 \times 47 \mu$Jy beam$^{-1}$. The beam is shown in the bottom left and a scale bar in the bottom right.

Figs 1(c) and (d) show the inner contours of VLA 3 coincide closely with the positions of MHO 745 A and B. K11 suggested the MHOs (745 A, B and C) are powered by a bipolar jet originating from VLA 2 aligned in the north–south direction. We measured the position angle of this jet to be between $-20^\circ$ and $-12^\circ$, where the former value was found by drawing a line between MHO 745 A and VLA 2 and the second was found between MHO 745 A and C (to the south-east of VLA 2).

The suggested jet is aligned within roughly $40^\circ$ of the large-scale outflow, which was found to have a position angle of $\sim 30^\circ$ (see Section 3.2.2). This is consistent with the classification criterion for an ionized jet used by (Purser et al. 2016), that the difference in deconvolved position angle between the jet and the outflow should be less than $45^\circ$. Therefore, the radio continua emission from VLA 2 and VLA 3 can be explained as a single compact source plus a jet elongated to the north in both 6 and 3.6-cm maps. However, due to its misalignment with the large-scale outflow, another explanation could be that VLA 3 is tracing the (shock-)ionized cavity walls of the large-scale outflow (see Section 4.1).

### 3.1.2 Other sources

The remaining detected sources are not known to be associated with IRAS 05373+2349. Following equation A2 of Anglada et al. (1998), which gives the number of background sources as a function of integrated flux based on data from Condon (1984), we expect approximately two background sources within the field of 4.5 by 5.5 arcmin at 6 cm, shown in Fig. 1(a).

VLA 1 is the first of two radio sources detected by Molinari et al. (2002). There is no infrared source found to be immediately associated with VLA 1 in any Spitzer IRAC (Fazio et al. 2004) or UKIDSS band. The source has a spectral index of $\alpha = -0.61 \pm 0.38$ which indicates synchrotron emission.

An angular size–frequency index, $\xi = -0.96 \pm 0.92$, was also derived for VLA 1. This index can be used to determine characteristics of the source. Yang et al. (2008) found that extragalactic jets show angular size–frequency indices of $\xi = -0.95 \pm 0.37$. The uncertainty in the measured angular size–frequency index is too large to draw such a conclusion. However, as there is no evidence of a YSO at this position, we find the source is likely extragalactic in origin.

The detailed analysis of the H$_2$ line emission by K11 did not focus as far afield as to where VLA 4–7 are located. Thus, it is difficult to associate any MHOs with these sources.

| Source name | $\lambda$ (cm) | RA (J2000) | Dec. (J2000) | Peak flux density (mJy beam$^{-1}$) | Integrated flux density (mJy) | Deconvolved size (arcsec) | P.A. ($^\circ$) | Convolved size (arcsec) |
|-------------|----------------|-------------|-------------|-----------------------------------|-------------------------------|--------------------------|--------------|--------------------------|
| VLA 1       | 6              | 05:40:21.31 | 23:52:09.80 | 0.22 $\pm$ 0.03                   | 0.40 $\pm$ 0.06               | 5.6 $\times$ 2.0          | 84 $\pm$ 173 | 9.2                      |
|             | 3.6            | 05:40:21.28 | 23:52:10.00 | 0.19 $\pm$ 0.02                   | 0.29 $\pm$ 0.03               | 2.5 $\times$ 1.6          | 49 $\pm$ 20  | 5.6                      |
|             | 1.3            | ...         | ...         | ...                               | ...                          | ...                      | ...          | ...                      |
| VLA 2       | 6              | 05:40:24.22 | 23:50:54.80 | 0.42 $\pm$ 0.02                   | 0.45 $\pm$ 0.04               | 1.6 $\times$ 0.4          | 39 $\pm$ 124 | 8.5                      |
|             | 3.6*           | 05:40:24.23 | 23:50:54.83 | 0.52 $\pm$ 0.03                   | 0.85 $\pm$ 0.08               | ...                      | ...          | ...                      |
|             | 1.3*           | 05:40:24.22 | 23:50:54.76 | 0.46 $\pm$ 0.03                   | 0.57 $\pm$ 0.12               | ...                      | ...          | ...                      |
| VLA 3       | 6              | 05:40:24.16 | 23:51:01.41 | 0.13 $\pm$ 0.02                   | 0.15 $\pm$ 0.02               | $< 2.1$                  | ...          | ...                      |
|             | 3.6            | 05:40:24.10 | 23:51:02.51 | 0.11 $\pm$ 0.02                   | 0.19 $\pm$ 0.03               | 2.9 $\times$ 1.9          | 15 $\pm$ 19  | 5.3                      |
|             | 1.3            | ...         | ...         | ...                               | ...                          | ...                      | ...          | ...                      |
| VLA 4       | 6              | 05:40:08.48 | 23:51:18.01 | 0.30 $\pm$ 0.04                   | 0.72 $\pm$ 0.13               | $< 1.3$                  | ...          | 14.2                     |
|             | 3.6            | ...         | ...         | ...                               | ...                          | ...                      | ...          | ...                      |
|             | 1.3            | ...         | ...         | ...                               | ...                          | ...                      | ...          | ...                      |
| VLA 5       | 6              | 05:40:10.52 | 23:49:32.75 | 0.93 $\pm$ 0.03                   | 0.93 $\pm$ 0.07               | $< 0.8$                  | ...          | 8.6                      |
|             | 3.6            | 05:40:10.49 | 23:49:32.57 | 0.57 $\pm$ 0.03                   | 0.76 $\pm$ 0.07               | 1.8 $\times$ 1.6          | 101 $\pm$ 168 | 4.5                      |
|             | 1.3            | ...         | ...         | ...                               | ...                          | ...                      | ...          | ...                      |
| VLA 6       | 6              | 05:40:13.74 | 23:49:10.68 | 0.34 $\pm$ 0.02                   | 0.34 $\pm$ 0.02               | $< 1.3$                  | ...          | 6.3                      |
|             | 3.6            | 05:40:13.74 | 23:49:09.54 | 0.26 $\pm$ 0.02                   | 0.32 $\pm$ 0.03               | 1.9 $\times$ 0.6          | 24 $\pm$ 37  | 3.5                      |
|             | 1.3            | ...         | ...         | ...                               | ...                          | ...                      | ...          | ...                      |
| VLA 7       | 6              | 05:40:08.02 | 23:47:22.78 | 2.06 $\pm$ 0.06                   | 2.28 $\pm$ 0.18               | $< 0.5$                  | ...          | 10.4                     |
|             | 3.6            | 05:40:07.97 | 23:47:22.22 | ...                               | ...                          | ...                      | ...          | ...                      |
|             | 1.3            | ...         | ...         | ...                               | ...                          | ...                      | ...          | ...                      |

Notes. * indicates sources that were measured in an aperture defined by the 1σ contour, as they were not well-fitted by a Gaussian. An ellipsis (...) in the peak position columns indicate non-detections at that wavelength. A double dash indicates the property was unable to be determined. The 3.6 cm flux of VLA 7 could not be accurately determined as it was not possible to apply a correction for the primary beam response at this position.
VLA 4 was found to be clearly associated with IRAC emission at 3.6, 4.5, 5.8 and 7.9 μm and the K-band of the UKIDSS survey. A spectral index could not be determined as it was only detected at 6 cm in our observations. However, a non-detection at 3.6 cm of VLA 4 indicates this source likely has a negative spectral index.

VLA 5 was found to be associated with faint IRAC emission at 3.6 and 4.5 μm. A spectral index of \( \alpha = -0.38 \pm 0.24 \) was found, consistent with a non-thermal source or narrowly consistent with optically-thin free–free emission.

There is faint IRAC emission associated with VLA 6 at 3.6 and 4.5 μm, which has a spectral index of \( \alpha = -0.15 \pm 0.28 \). It could equally be associated with a non-thermal or optically thin thermal source due to the uncertainty in the spectral index.

VLA 7 was found to be associated with faint IRAC emission at 3.6, 4.5 and 5.8 μm. The 3.6-cm flux of VLA 7 could not be accurately determined. This was due to the primary beam response not being accurately known so far from the pointing centre of the observation. However, it can be seen from our observations that the 3.6-cm detection of VLA 7 is much weaker than that at 6 cm, and therefore it likely has a negative spectral index.

Further observations of these sources will be required to determine their association with the IRAS 05373+2349 cluster or whether they are unrelated or extragalactic objects.

### 3.2 CARMA \(^{12}\)CO (J=1–0) and 2.7-mm continuum

In this section the \(^{12}\)CO (J=1–0) line and millimetre continuum observations taken towards VLA 2 are used to infer the properties of the dense molecular gas surrounding the region.

#### 3.2.1 2.7-mm continuum

Continuum emission at 2.7 mm was only detected towards VLA 2. Fig. 3 shows the continuum emission from the region around VLA 2 at 111.1 GHz (~2.7 mm) with MHOs 745 A, B and C (K11) indicated as crosses. The morphology of the emission shows a compact core with an extension from the south of the source in the south-east direction. The integrated and peak flux densities, measured in a 1σ contour, were 53.7 ± 2.6 mJy and 30.1 ± 1.6 mJy, respectively. The peak position, measured at the brightest pixel, was 05\(^{h}\)40\(^{m}\)24\(^{s}\)18.1 +23\(^{\circ}\)50\(^{\prime}\)55.01 (J2000).

A 2D-Gaussian was also fitted to the source. However, we found that only the inner compact core could produce a satisfactory fit; we derived an integrated flux density of 53.7 ± 5.9 mJy, a peak flux density of 29.9 ± 3.3 mJy and a deconvolved source size of 4.1 ± 0.4 by 2.9 ± 0.6 arcsec and a P.A. of 162 ± 18°. The deconvolved size corresponds to 4900 × 3500 au at 1.2 kpc. The peak position was 05\(^{h}\)40\(^{m}\)24\(^{s}\)194.2 +23\(^{\circ}\)50\(^{\prime}\)55.346 (J2000), with a positional uncertainty on the order of 0.5 arcsec (Ikarashi et al. 2011). Therefore the 2.7-mm source is associated with VLA 2.

To investigate how much millimetre emission might arise from the ionized material, we extrapolated the combined 6-cm flux of VLA 2 and VLA 3 to 2.7 mm, assuming a spectral index of \( \alpha = 0.38 \pm 0.14 \) and that the spectrum did not turn over. We found an expected flux density of \( \sim 1.95 \pm 0.73 \) mJy, which is \( \sim 4 \) per cent of the measured millimetre continuum (\( \sim 54 \) mJy). Therefore, the contribution of ionized gas emission in this case is negligible. However, if we use \( \alpha = 1.20 \pm 0.24 \), determined by excluding the 1.3 cm fluxes, the estimated contribution to the combined integrated flux of VLA 2 and 3 at 2.7 mm becomes \( \sim 25 \) mJy, roughly half of the measured millimetre continuum. Scaling the 6 cm fluxes of VLA 2 and 3 by \( \alpha = 1.2 \) and \( \alpha = 0.45 \), respectively, where the latter is the measured spectral index for VLA 3, gives a total 2.7 mm flux of \( \sim 20 \) mJy. We calculated the core mass for both cases below.

The mass of the material surrounding VLA 2, traced by the millimetre continuum emission was calculated using

\[
M_{\text{dust}} = \frac{d^3 S_v}{B_v(T_D) \kappa_v} \tag{3}
\]

where \( d = 1.2 \) kpc and \( T_D = 27 \) K are the distance and dust temperature (Molinari et al. 2000). \( S_v \) is the integrated flux density measured in the 1σ contour at \( \lambda = 2.7 \) mm. \( B_v(T_D) \) is the value of the Planck function at \( T_D \). Finally the dust opacity, \( \kappa_v = 0.1807 \) cm\(^2\) g\(^{-1}\), was found by extrapolating model opacity data (Ossenkopf & Henning 1994) to 2.7 mm, assuming a gas density of \( 1 \times 10^3 \) g cm\(^{-3}\) and thin ice mantles, where \( \beta \), the dust opacity index with frequency, was fitted to be \(-1.82\). A gas-to-dust ratio of 100 was assumed to infer the total core mass. In the case of all of the 2.7 mm emission originating from the dust, the total core mass is \( 23 \) M\(_{\odot}\). This is in close agreement with the value of \( \sim 26 \) M\(_{\odot}\) found by Molinari et al. (2008) from fitting the SED of the source. In the case of 25 mJy of the 2.7 mm emission originating from the ionized gas, and 29 mJy from the dust, we determined the core mass to be \( 12 \) M\(_{\odot}\).

#### 3.2.2 \(^{12}\)CO (J=1–0) line emission

Fig. 4 presents the \(^{12}\)CO (J=1–0) line profile of the large-scale outflow from VLA 2. High-velocity line wings are present in the spectrum. The inner, low velocity channels show evidence for self-absorption and/or missing flux.

Fig. 5 shows blueshifted and redshifted \(^{12}\)CO (J=1–0) emission integrated over several velocity ranges along with the intensity weighted first moment map of the line emission shown in grey-scale. Fig. 5 also shows six additional markers to that shown in Fig. 3. The two crosses about 40 arcsec to the north-east of VLA 2 show the positions of MHOs 741 A and B detected by K11, and the white star shows the position of a Class II pre-main-sequence star reported by Gutermuth et al. (2009). The three remaining markers to the south-west of VLA 2 show the positions of MHOs 738 A, B and C from K11, indicated as crosses.
The \(^{12}\)CO \((J=1-0)\) emission appears to show a bipolar outflow centred on VLA 2, appearing in the same orientation as previously seen by Zhang et al. (2005). There is also blueshifted and redshifted emission \(\sim 35\) arcsec to the north-east of VLA 2 that appears to resemble another resolved bipolar outflow. However, the only apparent candidate driving source of this outflow is a Class II source reported by Gutermuth et al. (2009), which does not lie directly between the red and blue lobes of the candidate outflow. Therefore, it is unlikely that this source is driving the outflow and we conclude that the \(^{12}\)CO \((J=1-0)\) emission in fact belongs to the single outflow centred on VLA 2. This could be due to inclination effects. For instance, if the outflow is almost in the plane of the sky, part of the blueshifted cavity will appear redshifted.

The groups of MHOs to the north-east and south-west are likely associated with the outflow from VLA 2, as they lie at opposite ends of the outflow seen in \(^{12}\)CO \((J=1-0)\) emission.

We also detected redshifted emission extending to the south-east of VLA 2 which ends in a patch of emission at \(05^h 40^m 26.7 + 23^\circ 50' 26''\) (J2000). However, there are no other signposts of star formation at that position.

### 3.2.3 Outflow properties

Table 4 presents the derived properties for the redshifted and blueshifted lobes of the outflow centred on VLA 2. The rows from top-to-bottom show: velocity range, mass, momentum, kinetic energy, mean velocity, length, dynamical time-scale, mass transfer rate, mechanical force and mechanical luminosity for the blue and red lobes of the outflow.

The velocity range shows the range of velocity channels that were integrated to produce the outflow properties. The inner, low velocity channels were excluded due to self-absorption or missing flux; channels with insignificant emission (< \(3\sigma\)) in the high-velocity wings were also excluded.

The outflow mass, \(M_{\text{outflow}}\), was derived for each lobe following Scoville et al. (1986):

\[
M(M_\odot) = 2.29 \times 10^{-5} \frac{(T + 0.926)}{e^{-5.55/T}} \frac{\tau}{e^{-\tau}} d_{\text{kpc}}^2 \int S_\nu \, dv
\]

where \(\tau\) is the optical depth, which we set to zero, assuming the source to be optically thin, and therefore the third term of equation (4) was set to unity. The distance, \(d_{\text{kpc}}\), was assumed to be 1.2 kpc. \(S_\nu\) is the integrated flux density in velocity channel, \(\nu\), that was found by summing up the emission within a region defined by the 1\(\sigma\) = 0.5 Jy beam\(^{-1}\) contours (the maximal noise level within the spectra). \(T\) is the excitation temperature of the outflow gas, which we assumed to be equal to the full Planck function derived brightness temperature, given as:

\[
T_B(K) = 0.048 [\nu \text{ (GHz)}] \times \ln \left( 1 + \frac{3.92 \times 10^{-8} [\nu \text{ (GHz)}] \theta''^2}{F_\nu \text{ (Jy beam}^{-1})} \right)^{-1}
\]

where \(\theta\) is the beamsize and \(F_\nu\) is the peak flux density of the outflow at frequency \(\nu\). For the derived properties listed in Table 4.
the largest peak flux from either blueshifted or redshifted lobe was used to define the brightness temperature for that outflow, which we found to be 6.47 Jy beam$^{-1}$ at 5.3 km s$^{-1}$ from the redshifted lobe. We determined the brightness temperature to be 27.3 K and found the total mass contained within the flows to be $\sim$0.8 M$_\odot$.

To calculate the outflow momentum, $P$, and kinetic energy, $E$, the $\int S_v \, dv$ term in equation (4) was replaced by $\int S_v \, (v - v_{\text{LSR}}) \cos i \, dv$ for the momentum, and $\int S_v \, [(v - v_{\text{LSR}}) / \cos i] \, dv$ for the kinetic energy, where $v$ is the velocity of each channel, $v_{\text{LSR}} = 2.3$ km s$^{-1}$ is the systemic velocity of the cloud taken from Zhang et al. (2005) and $i$ is the inclination of the outflow.

The columns in Table 4 show the properties uncorrected for inclination ($i = 0^\circ$) and the values in parentheses show the properties corrected using a mean inclination angle of $i = 57.3\degree$ (Bontemps et al. 1996). The uncorrected results are more closely relatable to the results of Zhang et al. (2005), who also assumed an inclination of 0$^\circ$. Although we have not applied this, a further correction of 3.5 may be applied to correct for optical depth effects (Bontemps et al. 1996).

The mass weighted velocity, $v_{\text{outflow}}$, of the outflow is given by the $P/M_{\text{outflow}}$ of the flow. The length of flow, $L$, was measured from the position of VLA 2 to the furthest $3\sigma$ contour (see Fig. 5) and was corrected for inclination.

From this we define the dynamical time-scale, $t_{\text{dyn}} = L/v_{\text{outflow}}$, and compute: the mass transfer rate, $M = M_{\text{outflow}}/t_{\text{dyn}}$; the mechanical force, $F = P/t_{\text{dyn}}$; and the mechanical energy transport rate, $E$, or mechanical luminosity, $L$, defined as $E/t_{\text{dyn}}$.

4 DISCUSSION

In addition to spectral analysis, a comparison of the radio-to-bolometric luminosity of a source can be used as a means of source classification. For low luminosity objects to have a strong radio flux, the emission can only arise from shock-ionized gas, as there is negligible emission via photoionization of circumstellar material. At higher luminosities, the expected Lyman continuum flux would dominate the observed radio flux, if the emission were produced via photoionization i.e. by an H ii region.

Fig. 6 presents the 6-cm radio flux of VLA 2 ($S_v = 0.45$ mJy) at a bolometric luminosity of 430 L$_\odot$, shown as a (blue) hexagon. The bolometric luminosity was found by taking an average of the luminosities found in the literature: 1100 L$_\odot$ at a distance of 2.0 kpc (found by the RMS survey, Lumsden et al. 2013), which scales to a luminosity of 396 L$_\odot$ at $d = 1.2$ kpc; 470 L$_\odot$ (Molinari et al. 2000), 400 L$_\odot$ (Molinari et al. 2008), 600 L$_\odot$ (Zhang et al. 2005) and 290 L$_\odot$ (K11); all at a distance of 1.2 kpc.

Also plotted on Fig. 6 are examples of jet-like objects taken from Anglada (1995) shown as green crosses, YSOs and H ii-regions detected in the Red MSX Source (RMS) survey (Lumsden et al. 2013) shown as red circles and black plus signs, respectively. At the average bolometric luminosity of 430 L$_\odot$, VLA 2 appears to be associated with the jet-like objects.

The luminosity of 290 L$_\odot$, that was used in our average, was found via SED modelling of VLA 2 (K11). The authors used a large aperture ($\sim 22 000$ au) and incorporated fluxes from lower resolution mid-infrared to radio wavelengths in the fitting, which suggested a Class 0/I actively accreting protostar with a mass of about 4.5 M$_\odot$. A separate SED fitting yielded a luminosity of 6030 L$_\odot$ ($d = 1.12$ kpc). However, this is likely incorrect as the authors modelled the system as edge on, which greatly increased the fitted luminosity. Thus, this value was not included in our average.

4.1 Small-scale jet

Multiple shock-features were detected in the region around VLA 2 in H$_2$ line emission (K11), also see Varricatt et al. 2010, extending to the north and south, as well as in the north-east south-west directions, roughly aligning with the large-scale outflow we observe in $^{13}$CO (J=1–0).

Fig. 7 presents the positions of MHOs 745 A, B and C, suggested a Class 0/I actively accreting protostar with a mass of about 4.5 M$_\odot$. This suggests a source with varying spectral index or, alternatively, two discrete objects with one being detected only at 6 cm and the other detected only at 3.6 cm. The MHOs also coincide with the blue- and redshifted $^{12}$CO (J=1–0) emission (see Fig. 5). This suggests that VLA 2 and the MHOs could also arise from shocks created by jet material impinging upon the cavity walls created by the large-scale outflow.

4.2 Large-scale outflow

Our $^{12}$CO (J=1–0) observations further resolved the previously observed large-scale outflow (Zhang et al. 2005) centred on VLA 2. The derived properties for the outflow (listed in Table 4) are lower than those found by the single dish observations of Zhang et al. (2005), which were $M_{\text{outflow}} = 3.4$ M$_\odot$, $P = 27.1$ M$_\odot$ km s$^{-1}$ and $E = 34 \times 10^{42}$ erg. However, this is expected due to the fraction of flux that is recovered by interferometry compared to single dish
observations, which we estimate to be around 10 per cent for the line centre (e.g. Qiu et al. 2009; Kwon et al. 2015). Interferometric observations of low-mass outflows can expect to find outflow masses on the order of $10^{-3}$–$10^{-2}$ M⊙ (e.g. Tamura et al. 1996). Single dish observations of massive outflows find outflow masses in the range $\sim$5–1000 M⊙ (Maud et al. 2015). Assuming that 10 per cent of the flux in the single dish observations is recovered by the interferometer, we therefore expect massive YSOs to have interferometric outflow masses of 0.5–100 M⊙.

Therefore, the outflow detected in 12CO (J=1–0), found to have a total mass of $\sim$0.8 M⊙, is higher than that observed for low-mass YSOs and is consistent with the lower end of the expected outflow masses found by Maud et al. (2015).

5 CONCLUSIONS

We present observations of the embedded cluster associated with IRAS 05373+2349 made with the VLA in its C configuration at wavelengths of 6, 3.6 and 1.3-cm and CARMA observations of 12CO(J=1–0) line emission and 2.7-mm continuum. The spectral index of VLA 2 derived from the centimetre observations was found to be consistent with an ionized wind. In addition, the integrated radio flux of VLA 2 compared to the average bolometric luminosity of the associated IRAS source indicated a jet-like object.

We also detected a new radio source, VLA 3, that is coincident with MHOs previously observed in 12CO, line emission around VLA 2. This supports the claim by K11 that the MHOs trace a bipolar jet centred on VLA 2, which is roughly aligned with the large-scale outflow presented here in 12CO(J=1–0). However, we did not find evidence of any of the other outflows in the region, as proposed by K11. Therefore, we suggest that the large-scale outflow likely arises from a single jet centred on VLA 2 and traced by VLA 3. Furthermore, the difference in position angle of $\sim$40–50° between the jet and the large-scale outflow could be explained by precession of the jet. However, an alternative explanation of VLA 3 could be (shock-)ionization of the outflow cavity walls of the large-scale outflow.

Finally, our observations of the millimetre continuum and 12CO (J=1–0) emission towards VLA 2 found a core mass of between 12 and 23 M⊙ and an outflow mass of 0.8 M⊙, intermediate between outflows from low- and high-mass YSOs.

The goal of this study was to investigate whether the properties of the IM protostar IRAS 05373+2349 VLA 2 are similar to that of low- and/or high-mass protostars. Our observations of the radio continuum emission from IRAS 05373+2349 VLA 2 indicate that it arises in a jet, similar to those observed towards low- and high-mass young stars (e.g. Anglada 1995; Purser et al. 2016). In addition, we detect radio continuum emission (VLA 3) associated with previously observed H2 line emission which would indicate shock-ionized knots in the jet from this source, seen commonly in the sample of jets from high-mass YSOs observed by Purser et al. (2016). The core mass lies between those observed for low- and high-mass YSOs, with the youngest low-mass sources having core masses on the order of a solar mass (Stutz et al. 2013), and the mass reservoirs of high-mass stars reaching thousands of stellar masses (Beuther et al. 2002). The jet itself is accompanied by an outflow from VLA 2, seen in 12CO (J=1–0) emission, which has properties intermediate between those seen towards low- and high-mass YSOs.

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REFERENCES

Anglada G., 1995, Rev. Mex. Astron. Astrofís. Ser. Conf., 1, 67
Anglada G., Villuendas E., Estellera R., Beltrán M. T., Rodríguez L. F., Torrelles J. M., Curiel S., 1998, AJ, 116, 2953
Beltrán M. T., 2015, Ap&S, 355, 283
Beltrán M. T., de Wit W. J., 2016, A&AR, 24, 6
Beltrán M. T., Girart J. M., Estellera R., Ho P. T. P., Palau A., 2002, ApJ, 573, 246
Beltrán M. T., Girart J. M., Estellera R., Ho P. T. P., Anglada G., 2008, A&A, 481, 93
Beuther H., Schilke P., Menten K. M., Motte F., Sridharan T. K., Wyrowski F., 2002, ApJ, 566, 945
Beuther H., Churchwell E. B., McKee C. F., Tan J. C., 2007, in Reipurth B., Jewitt D., Keil K., eds, Protostars and Planets V. Univ. Arizona Press, Tucson, p. 165
Bontemps S., André P., Terebey S., Cabrit S., 1996, A&A, 311, 858
Condon J. J., 1984, ApJ, 287, 461
Crimier N. et al., 2010, A&A, 516, A102
Fazio G. G. et al., 2004, ApJS, 154, 10
Fontani F., Zhang Q., Caselli P., Bourke T. L., 2009, A&A, 499, 233
Fuente A., Neri R., Martín-Pintado J., Bachiller R., Rodriguez-Franco A., Palla F., 2001, A&A, 366, 873
Fuente A., Ceccarelli C., Neri R., Alonso-Albi T., Caselli P., Johnstone D., van Dishoeck E. F., Wyrowski F., 2007, A&A, 468, L37
Fuente A. et al., 2009, A&A, 507, 1475

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