SDSS J163459.82+204936.0: A RINGED INFRARED-LUMINOUS QUASAR WITH OUTFLOWS
IN BOTH ABSORPTION AND EMISSION LINES

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ABSTRACT

SDSS J163459.82+204936.0 is a local (z = 0.1293) infrared-luminous quasar with LIR = 10^{11.91} L_☉. We present a detailed multiwavelength study of both the host galaxy and the nucleus. The host galaxy, appearing as an early-type galaxy in the optical images and spectra, demonstrates violent, obscured star formation activities with SFR ≈ 140 M_☉ yr⁻¹, estimated from both the polymeric aromatic hydrocarbon emission or IR luminosity. The optical to NIR spectra exhibit a blueshifted narrow cuspy component in Hβ, HeⅠλ5876, 10830, and other emission lines consistently with an offset velocity of ∼900 km s⁻¹, as well as additional blueshifting phenomena in high-ionization lines (e.g., a blueshifted broad component of HeⅠλ10830 and the bulk blueshifting of [OⅢ]5007). While there exist blueshifted broad absorption lines (BALs) in NaⅠ D and HeⅠλ3889, 10830, indicative of the active galactic nucleus outflows producing BALs and emission lines. Constrained mutually by the several BALs in the photoionization simulations with Cloudy, the physical properties of the absorption line outflow are derived as follows: density 10⁴ < n_H < 10⁵ cm⁻³, ionization parameter 10⁻¹³ < ζ < 10⁻⁰⁷, and column density 10²².⁵ < N_H < 10²⁹.⁹ cm⁻², which are similar to those derived for the emission line outflows. This similarity suggests a common origin. Taking advantages of both the absorption lines and outflowing emission lines, we find that the outflow gas is located at a distance of ∼48–65 pc from the nucleus and that the kinetic luminosity of the outflow is 10⁴⁴–10⁴⁶ erg s⁻¹. J1634+2049 has a off-centered galactic ring on the scale of ∼30 kpc that is proved to be formed by a recent head-on collision by a nearby galaxy for which we spectroscopically measure the redshift. Thus, this quasar is a valuable object in the transitional stage emerging out of dust enshrouding as depicted by the co-evolution scenario invoking galaxy merger (or violent interaction) and quasar feedback. Its proximity enables our further observational investigations in detail (or tests) of the co-evolution paradigm.

Key words: galaxies: active – galaxies: individual (SDSS J163459.82+204936.0) – galaxies: starburst – quasars: absorption lines – quasars: emission lines

1. INTRODUCTION

In the generally believed dark matter (CDM) paradigm of the universe, galaxies grow in a “bottom-up” fashion as led by the CDM halos, with smaller ones forming first and then merging into successively larger ones. Mergers and strong interactions of gas-rich galaxies are also believed to play a vital role in triggering the accretion activity, namely the active galactic nucleus (AGN) phenomenon, of supermassive black holes (SMBHs) which reside at the centers of most (if not all) galaxies (see, e.g., Hopkins et al. 2006). Observationally, the ultraluminous infrared galaxies (ULIRGs; LIR > 10¹² L_☉) in the local universe that were discovered three decades ago are found mostly in mergers, which inspired a merger-driven evolutionary sequence from ULIRGs to quasars and finally to present-day elliptical galaxies (Sanders et al. 1988; Sanders & Mirabel 1996; Hopkins et al. 2006, 2008). At first, galaxy merging induces enormous starbursts, which are almost completely enshrouded by dust (i.e., in the ULIRG phase), and triggers the central AGN; with the increasing feedback from the starbursts and AGN, the cold gas and dust are heated up and even expelled out of the galaxy and thus the AGN becomes optically bright (i.e., the quasar phase); meanwhile, the large-scale starbursts decline. Finally, the cold gas and dust is gone, the AGN shuts down, and the galaxy becomes an old elliptical. This scenario has been supported by subsequent observations and N-body/SPH simulations (Hopkins et al. 2006). Particularly, the tight correlation between the masses of SMBHs and the properties of the spheroids observed in local quiescent galaxies suggests a co-evolution of galaxies with SMBHs (see Kormendy & Ho 2013 for a review).

In practice, however, the concrete triggering and feedback processes underlying this scenario have remained unknown for the last decades. This leaves many open questions, e.g., what the timing is between starburst and AGN activities, how the AGN feedback operates in the host galaxy. The physical processes actually cannot be learned from the statistical studies alone (e.g., correlation analysis) of galaxy and quasar samples; they are also beyond the capabilities of current simulations (see, e.g., Hopkins et al. 2006; Veilleux et al. 2009). A straightforward way is to carry out complementary...
investigations of individual sources in detail, particularly of the rare cases in transitional phases of the proposed evolutionary sequence.

This paper presents a detailed multiwavelength analysis of SDSS J163459.82+204936.0 (hereafter J1634+2049), a type-1 AGN at $z = 0.1293$ with outflows revealed in both broad absorption lines (BALs) and narrow emission lines. This object was detected by the Infrared Astronomical Satellite (IRAS) and was compiled by Condon et al. (1995) into their catalog of radio-detected bright IRAS sources. J1634+2049 was noted by us from the SDSS spectral data set when we compiled the sample of low-$z$ quasars with broad He$\lambda$3889 absorption troughs (Liu et al. 2015). In the present paper, we will see that J1634+2049 is a LIRG with a total infrared luminosity of $10^{11.91} L_\odot$, suggesting a strong ongoing star formation ($\text{SFR} \sim 140 M_\odot \text{yr}^{-1}$). Both star formation regions and the AGN show considerable internal dust obscuration. The spectroscopy observation on the nearby galaxy demonstrates that J1634+2049 was collided through by a galaxy, leaving a stellar ring around it on scales of 30 kpc (Section 3.3). Besides the outflow revealed in BALs of He$^+$ and Na$\lambda$ D (Section 2.4.3), there are outflows revealed in the emission of the narrow Hydrogen Balmer and Paschen lines, [O$\Pi$], [O$\Pi$] $\lambda$3507, He$\lambda$5876, and H1030 (Section 2.4.2). Analyses with model calculations indicate that the physical conditions of the absorption line outflow and emission line outflow are similar, suggesting that the two are intrinsically the same outflow. The outflow is estimated to be $\sim$48–65 pc away from the central nuclei with a large kinetic luminosity $10^{44} \text{erg s}^{-1}$ (see Section 3.4). In terms of the mid-infrared (MIR)—far-infrared (FIR) spectral energy distribution (SED) (see Section 2.2), J1634+2049 is between the prototypical ULIRG/quasar composite object Mrk 231 (see e.g., Kawakatu et al. 2006; Veilleux et al. 2013; Spoon et al. 2013; Leightly et al. 2014) and normal quasars. Mrk 231 has been long-known as the nearest ULIRG/QSO composite object, with a total infrared luminosity of $3.6 \times 10^{12} L_\odot$, an AGN bolometric luminosity of $1.5 \times 10^{46} \text{erg s}^{-1}$, and a star formation rate (SFR) of $170 M_\odot \text{yr}^{-1}$ (Veilleux et al. 2013). It is a FeLoBAL quasar and displays neutral and ionized nuclear outflows in several optical and UV tracers (e.g., Veilleux et al. 2013; Leightly et al. 2014); recently Mrk 231 has been regarded as the archetype showing galactic-scale quasar-driven winds. J1634+2049 should be a young, transitional quasar immediately after Mrk 231 in the evolutionary sequence, blowing out of the enshrouded dust after a violent galactic collision. Throughout this work we assume a cosmology with $H_0 = 70 \text{km s}^{-1} \text{Mpc}^{-1}$, $\Omega_m = 0.3$, and $\Omega_\Lambda = 0.7$.

2. OBSERVATIONS AND DATA ANALYSIS

2.1. Spectroscopic and Photometric Observations

J1634+2049 has been observed in multiple bands both spectroscopically and photometrically. It was first spectroscopically observed by SDSS on 2004 August 7 UT with an exposure time of 3072 s under the seeing of $\sim 1^\prime 3$, with a wavelength coverage of 3800–9200 Å. The SDSS pipeline gave a redshift of $0.1286 \pm 0.0014$. We measure a redshift $z_{\text{em}} = 0.1293 \pm 0.0007$ from [S$\Pi$] $\lambda\lambda$6716, 6731, and all the following rest frame spectra are referred to this redshift.

Trying to extend the optical spectrum toward both the near-ultraviolet and near-infrared ends, we have taken spectroscopy with the Double Spectrograph (DBSP) on the Palomar 5 m Hale telescope (Oke & Gunn 1982). Two exposures of 300 s each were obtained on 2014 April 23 UT, when the sky was basically clear and the seeing was $\sim 1^\prime 5$. With a $1^\prime 5$ slit-width and the 600/4000 grating, the blue side (3150–5700 Å) spectrum has a spectral resolution of $\sim 4.13 \AA$; with a $1^\prime 5$ slit-width and the 600/10000 grating, the red-side (7800–10200 Å) spectrum has a spectral resolution of $\sim 4.2 \AA$. The data reduction was performed with the standard routines in the IRAF.

In addition, there are two small nearby galaxies seen to the west of J1634+2049 in the SDSS image (C1, C2 in Figure 2). The two galaxies show similar colors to J1634+2049 and their photometric redshift values given by SDSS are 0.224 ± 0.0421 and 0.283 ± 0.0781, respectively, close to that of J1634+2049 in light of the large uncertainties of the photometric redshifts. Considering the galactic ring around J1634+2049, we wonder if the two nearby galaxies had been once interacting with J1634+2049. To measure the redshifts of two possible physically companion galaxies, we performed spectroscopy observations of J1634+2049 and the two nearby galaxies using the Yunnan Faint Object Spectrograph and Camera (YFOSC) mounted on the Lijiang GMG 2.4 m telescope on 2015 March 13. The G10 (150 mm$^{-1}$) grating provides a wavelength range of 3400–10000 Å and a resolution of $R \approx 760$. The $1^\prime 8$ slit-width was adopted and the slit was rotated by a position angle PA = 86° to place J1634+2049 and C1 and C2 (see Figure 9) all in the slit. Two exposures of 2400 s each were obtained. The data reduction was performed with the standard IRAF routines. Due to the low spectral resolution and the imperfect HeNeAr lamp spectra, IRAF gives a quite large uncertainty of the wavelength calibration, 2.84 Å (rms).

The near-infrared (NIR) spectroscopic observations for this object were performed with the TripleSpec spectrograph on the Palomar 5 m Hale telescope on 2012 April 15. Four exposures of 120 s each were obtained in an A-B-B-A dithering mode, and the sky was clear with seeing $\sim 1^\prime 2$. The slit-width of TripleSpec was fixed to $1^\prime$. Two telluric standard stars were taken quasi-simultaneously. The data was reduced with the IDL program SpexTool (Cushing et al. 2004). The flux calibration and telluric correction were performed with the IDL program using the methods described in Vacca et al. (2003).

MIR Observation of J1634+2049 was performed using the Infrared Spectrograph (IRS; Houck et al. 2004) on board Spitzer (Werner et al. 2004) on 2008 April 30 (PI: Lei Hao, program ID: 40991). All four modes—short low 1 (SL1), short low 2 (SL2), long-low 1 (LL1), and long-low 2 (LL2)—were used to obtain full 5–55 μm low-resolution ($R \sim 100$) spectra. We obtain the reduced spectrum from the public archive “Cornell Atlas of Spitzer/IRS Sources”$^9$ (CASSIS v7; Lebouteiller et al. 2011).

J1634+2049 has been photometrically observed in multiple bands and we list all the available photometric data in Table 1.

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$^9$ IRAF is distributed by the National Optical Astronomy Observatories, which are operated by the Association of Universities for Research in Astronomy, Inc., under cooperative agreement with the National Science Foundation.

$^10$ The Cornell Atlas of Spitzer/IRS Sources (CASSIS) is a product of the Infrared Science Center at Cornell University, supported by NASA and JPL. http://cassis.astro.cornell.edu/atlas/.
Table 1
Photometric Data

| Band     | Mag/Flux  | Facility | Obs. Date   | Reference |
|----------|-----------|----------|-------------|-----------|
| FUV      | 21.75 ± 0.46 mag | GALEX    | 2006 Dec 23 | 1         |
| NUV      | 20.47 ± 0.17 mag | GALEX    | 2006 Dec 23 | 1         |
| Petrosian $u$ | 18.94 ± 0.04 mag | SDSS    | 2003 Jun 23 | 2, 3      |
| Petrosian $g$ | 17.76 ± 0.01 mag | SDSS    | 2003 Jun 23 | 2, 3      |
| Petrosian $r$ | 16.95 ± 0.01 mag | SDSS    | 2003 Jun 23 | 2, 3      |
| Petrosian $i$ | 16.30 ± 0.01 mag | SDSS    | 2003 Jun 23 | 2, 3      |
| Petrosian $z$ | 16.07 ± 0.01 mag | SDSS    | 2003 Jun 23 | 2, 3      |
| $J$      | 14.65 ± 0.04 mag | 2MASS   | 1997 Jun 09 | 4         |
| $H$      | 13.45 ± 0.03 mag | 2MASS   | 1997 Jun 09 | 4         |
| $K_s$    | 12.25 ± 0.03 mag | 2MASS   | 1997 Jun 09 | 4         |
| $W1$     | 10.73 ± 0.02 mag | WISE    | 2010 May 29 | 5         |
| $W2$     | 9.72 ± 0.02 mag | WISE    | 2010 May 29 | 5         |
| $W3$     | 7.07 ± 0.02 mag | WISE    | 2010 Feb 21 | 5         |
| $W4$     | 4.61 ± 0.02 mag | WISE    | 2010 Feb 21 | 5         |
| IRAC 8 $\mu$m$^a$ | 0.035 ± 0.001 Jy | Spitzer | 2008 Apr 30 | 6         |
| IRAS 12 $\mu$m | 0.085 ± 0.019 Jy | IRAS    | 1991 May 30 | 7         |
| IRAC 16 $\mu$m$^b$ | 0.067 ± 0.002 Jy | Spitzer | 2008 Apr 30 | 6         |
| IRS 22 $\mu$m$^a$ | 0.110 ± 0.003 Jy | Spitzer | 2008 Apr 30 | 6         |
| MIPS 24 $\mu$m$^a$ | 0.114 ± 0.003 Jy | Spitzer | 2008 Apr 30 | 6         |
| IRAS 25 $\mu$m | 0.141 ± 0.018 Jy | IRAS    | 1991 May 30 | 7         |
| IRAS 60 $\mu$m | 0.559 ± 0.045 Jy | IRAS    | 1991 May 30 | 7         |
| AKARI 65 $\mu$m | 0.239 Jy | AKARI  | 2011 Sep 08 | 8         |
| AKARI 90 $\mu$m | 0.579 ± 0.060 Jy | AKARI   | 2011 Sep 11 | 8         |
| IRAS 100 $\mu$m | 1.172 ± 0.199 Jy | IRAS    | 1991 May 30 | 7         |
| AKARI 140 $\mu$m | 1.468 ± 1.753 Jy | AKARI   | 2011 Sep 11 | 8         |
| 1.4 GHz | 21.97 ± 0.147 mJy | FIRST   | 1998 Oct 07 | 9         |

Note.
$^a$ Synthetic photometry data from Spitzer IRS spectrum given by Spitzer data archive.
$^b$ Referenced as.

References.
(1) Morrissey et al. 2007; (2) York et al. 2000; (3) Abazajian et al. 2009; (4) Skrutskie et al. 2006; (5) Wright et al. 2010; (6) Houck et al. 2004; (7) Moshir et al. 1992; (8) Doi et al. 2009; (9) Becker et al. 1994.

2.2. Broadband SED

As shown in Figure 1, we construct the broadband SED in rest frame wavelength using the photometric data and spectra in multiple bands. These data are corrected for Galactic extinction using the dust map of Schlegel et al. (1998) and the Fitzpatrick (1999) reddening curve. Because these photometric and spectroscopic observations are non-simultaneous, we first check the variability of this object before we analyze the broadband SED. The Catalina Sky Survey11 performs an extensive photometric monitor since 2005 April 9 (MJD from 53469 to 56590), and has 272 observations so far. We obtain these data from the Catalina Surveys Data Release 2 (CSDR2) and bin it every day. As the bottom panel of Figure 1 shows, J1634+2049 has a long-term optical variability within 0.2 mag in the V band. Such a variability amplitude does not impact our discussions on its SED and so on below.

The top panel of Figure 1 shows the broadband SED of J1634+2049. The average QSO spectrum from the UV to FIR band scaled at 2 $\mu$m is overplotted for comparison. This average QSO spectrum is combined from the UV to optical average QSO spectrum of Vanden Berk et al. (2001), the NIR average QSO spectrum of Glikman et al. (2006) and the FIR average QSO spectrum of Netzer et al. (2007). It is obvious that the observed SED of J1634+2049 shows a very different shape from that of the average QSO spectrum. In the UV, optical, and NIR $J$ and $H$ bands, J1634+2049 is much lower than the average QSO spectrum; from the K band up to the MIR ($5-30 \mu$m), the shape of the SED is similar to that of the average QSO spectrum; in the FIR band, J1634+2049 shows an obvious excess, 10 times higher than the luminosity of the average SED of QSOs at 60 $\mu$m.

Note that in Figure 1 the AKARI photometric flux densities are systematically lower than the IRAS ones, which is actually because the AKARI data at 65 and 140 $\mu$m are not reliable. For the AKARI data, the quality flags at 65 $\mu$m, 90 $\mu$m and 140 $\mu$m are “1,” “2,” and “3,” respectively, where “3” indicates the highest data quality and “1” indicates that the source is not confirmed. For the IRAS data, the quality flags of the flux densities at 12 $\mu$m, 25 $\mu$m, 60 $\mu$m, and 100 $\mu$m are “1,” “2,” “3,” and “4,” respectively, where “3” means the highest data quality and “1” means that the flux is only an upper limit. Thus, the IRAS flux densities at 25 $\mu$m, 60 $\mu$m, and the AKARI flux density at 90 $\mu$m are the most reliable; the IRAS flux density at 100 $\mu$m is the second most reliable with a quality flag of “2” and it is consistent with the AKARI flux densities at 90 $\mu$m within 1-$\sigma$. The IRAS 12 $\mu$m flux density is higher than the WISE W3 (12 $\mu$m), which is because the IRAS 12 $\mu$m datum is just an upper limit (with a IRAS quality flag of “1”). Besides, the flux density of the IRAS 25 $\mu$m is consistent with those of the WISE W4 and Spitzer/IRS 22 $\mu$m; the flux densities of

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11 The website is http://nessoi.cacr.caltech.edu/DataRelease/. The Catalina Sky Survey (CSS) is funded by the National Aeronautics and Space Administration under grant no. NNG05GF22G issued through the Science Mission Directorate Near-Earth Objects Programs. The CTRS survey is supported by the U.S. National Science Foundation under grant no. AST-0909182.
WISE W3 (12 μm) and W4 (22 μm) are consistent with those of SPITZER/IRAC 8 μm, SPITZER/IRAC 16 μm, and SPITZER/IRS 22 μm.

We try to match its SED to the reddened versions of the average QSO spectrum with different extinction curves. In Figure 1, the blue dashed line indicates the average QSO spectrum reddened with Milky Way extinction curve (Fitzpatrick 1999) by $E_{B-V} = 0.64$, while the purple and green dashed lines show the reddening with SMC extinction curve (Pei 1992) by $E_{B-V} = 0.61$ and the LMC extinction curve (Misselt et al. 1999) by $E_{B-V} = 0.66$, respectively. Here, the $R_V$ for the Milky Way extinction curve is 3.1, and for the LMC extinction curve is 2.6 (Weingartner & Draine 2001). It is hard to distinguish the extinction types according to the optical and NIR spectra, since these reddened average QSO spectra show little difference in the optical and NIR bands. Considering also the NUV and FUV photometric data retrieved from the GALEX archive (albeit not as superior as a UV spectrum), the LMC extinction curve is favored (see Figure 1). Hereafter, we will employ the LMC extinction curve with $R_V = 2.6$ for the internal dust obscuration of J1634+2049.

The IR luminosity (8–1000 μm) is calculated based on the IRAS photometric fluxes following Sanders & Mirabel (1996). Because the IRAS 12 μm flux density is an upper limit, in the calculation we use the WISE datum in the W3 band (12 μm) instead. It gives $\log L_{IR}(L_\odot) = 11.91 \pm 0.03$, which is very close to the defining IR luminosity of ULIRGs. As Schweitzer et al. (2006) suggested, for typical QSOs (e.g., PG QSOs), most of the far-infrared luminosity is originated from star formation. If this is the case for J1634+2049, following the relation $SFR (M_\odot \text{ yr}^{-1}) = 4.5 \times 10^{-44} L_{IR} (\text{erg s}^{-1})$ (Kennicutt 1998), the SFR is estimated to be $SFR = 140 \pm 43 M_\odot \text{ yr}^{-1}$. The scatter of this equation is $\pm 30\%$ (Kennicutt 1998), which is dominated the statistical uncertainty of the SFR. From the polycyclic aromatic hydrocarbon (PAH) emission lines and $24 \mu m$ continuum, the SFR are estimated to be $141 M_\odot \text{ yr}^{-1}$ and $143 M_\odot \text{ yr}^{-1}$ respectively, which are well consistent with the SFR estimated from its IR luminosity.

The $k$-corrected radio power at 1.4 GHz for J1634+2049 is also estimated, $P_{1.4 \text{GHz}} = 9.14 \times 10^{23} \text{ W Hz}^{-1}$. This is calculated by $P_{1.4 \text{GHz}} = 4\pi D_L^2 S_{1.4} / (1+z)^{1+\alpha_r}$, where the radio spectral index $\alpha_r (F_r \propto \nu^{\alpha_r})$ is assumed to be $-0.5$.

We compare the SED of J1634+2049 with that of Mrk 231, which is a prototypical nearest ULIRG/quasar composite object. The SED of Mrk 231 is constructed from the multiband spectra we collected. The FUV (1150–1450 Å) spectrum is obtained.

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**Figure 1.** Top: the broadband spectral energy distribution of J1634+2049 constructed from non-simultaneous photometry data (red) and spectra, corrected for Galactic reddening and brought to the rest frame. The black solid lines are the SDSS spectrum, the NIR spectrum observed by P200 TripleSpec, and the Spitzer IRS spectrum. Overplotted for comparison is the average QSO spectrum (gray dotted line; see the text); the reddened versions of the average QSO spectrum by different extinction curves are denoted by the orange (MW), purple (LMC), and green (SMC) dashed lines. The light blue dotted line is the SED of Mrk 231, constructed from the multiband spectra scaled at $\sim 2 \mu m$. Bottom: the variability of the V band of J1634+2049. The raw data (gray dots) are taken from Catalina Survey with the binned version within every day plotted also (black dot).
with \textit{Cosmic Origins Spectrograph} (COS) G130M grating on board the \textit{Hubble Space Telescope} (HST). The UV spectrum within (1600–3200 Å) is obtained with \textit{Faint Object Spectrograph} (FOS) G190 and G270 gratings on board HST. All the HST spectra are retrieved from HST data archive.\footnote{http://archive.stsci.edu/hst/search.php} The optical spectrum (3750–7950 Å) is obtained from Kim et al. (1995). We also observed Mrk 231 using TripleSpec spectrograph on Hale telescope on 2013 February 23 to obtain its NIR spectra. The FIR spectrum is obtained from Brauher et al. (2008). These multiband spectra are corrected for the Galactic extinction, and combined together by matching to its photometry data. In Figure 1 we scale the SED of Mrk 231 at 2 μm to that of J1634 +2049. As Figure 1 demonstrates, the shapes of the two SEDs differ significantly. Mrk 231 appears more reddened in the NUV–optical continuum. More strikingly, Mrk 231 has a much larger excess in MIR and FIR bands with a deeper silicate absorption dip at 9.7 μm than J1634 +2049.

2.3. Analysis of SDSS Images

J1634+2049 was photometrically observed by SDSS in the u, g, r, i, and z bands on 2003 June 23 UT, with an exposure time of 54 s per filter in drift-scan mode (Gunn et al. 1998). A bright point-like source appears at the center of its images and the whole galaxy is almost round in shape and has no spiral arms, indicating an elliptical/spheroidal or a face-on S0 galaxy. Closer inspection reveals a low surface brightness (SB) yet visible, circumgalactic ring-like structure. Although the standard SDSS images have a relatively short exposure time and low spatial resolution due to the seeing limit, they are still very useful for global study. The drift-scan mode, yielding accurate flat-fielding, in combination with the large field of view (FOV) ensures very good measurement of the sky background, and thus the azimuthally averaged, radial SB profile can be reliably determined down to μ_r \approx 27 mag arcsec^{-2} (e.g., Pohlen & Trujillo 2006; Erwin et al. 2008; Jiang et al. 2013). We try to perform a two-dimensional (2D) decomposition of the AGN and host galaxy of Mrk 231 using GALFIT (Peng et al. 2002, 2010). An accurate decomposition cannot only help us understand its host galaxy, but also put an independent constraint on the SED of the AGN, which is helpful for spectral fitting (cf. Jiang et al. 2013).

Taking advantage of the large FOV, a bright yet unsaturated nearby star (SDSSJ163503.44+204700.3) is selected as the PSF image, whose precision is essentially important to separate the AGN from the host. The host galaxy is represented by a Sérsic (1968) $r^{1/n}$ function. During the fitting, the sky background is set to be free; the outer ring region, enclosed by the green polygon as in Figure 2, is masked out. All other photometric objects in the field identified either by SExtractor (Bertin & Arnouts 1996) or by the SDSS photometric pipeline are also masked out.

We begin the fitting with a free PSF + Sérsic scheme allowing all parameters to vary, which yields an unreasonable high value of Sérsic index ($n > 10$). Then we thus try to fixed $n$ to 4, 3, 2, and 1, respectively (see Jiang et al. 2013). Except for the $u$-band image, which is totally dominated by the AGN component, all other four bands yield well consistent results: the best-fit Sérsic component is with $n = 4$. We have also attempted to add an exponential disk component, yet no convergent result can be achieved. The fitting results in the $g$, $r$, and $i$ bands are summarized in Figure 2 and Table 2.

2.4. Analysis of the Optical–NIR Spectrum

2.4.1. Decomposition of the Continuum

Before we perform the analysis on its spectra from different telescopes/instruments, we first check the aperture effect because J1634+2049 is an extended galaxy. Although the SDSS, DBSP, and NIR spectra are observed with different apertures/slits, we find that the three spectra, which were calibrated independently, are well consistent in flux level; especially, DBSP, and SDSS spectra are almost the same in their overlap part. On the other hand, as Figure 3 demonstrates, the surface brightness of J1634+2049 decreases rapidly in $r > 1''$ and therefore the outer region ($r > 1.5''$) has a negligible contribution to these spectra. Besides, considering that the SDSS and DBSP spectra are almost the same although the three spectra were observed in different time (SDSS: MJD 53224, NIR: MJD 56034, DBSP: MJD 56771), the variability effect can be ignored.

Figure 4 shows the spectrum combined from the DBSP, SDSS, and NIR spectra, the overlap parts of which are weighted by spectral S/N. First, we take a global overview of the features of the AGN and starlight components. In Section 2.3 the 2D imaging decomposition by GALFIT gives the contributions of the two components. Within the 30''-diameter aperture, the PSF (AGN) component accounts for 59% of the g-band light, 44% in the $r$, 44% in the $i$, and 67% in the $z$, respectively. The u-band image is much less sensitive, so an accurate decomposition is difficult; our rough decomposition shows that it is totally dominated by a PSF component. According to the 2D decomposition, the colors ($g - r$, $r - i$) of the Sérsic component are (0.93, 0.5), which suggest that the age of the dominating stellar component could be older than 10 Gyr (Bruzual & Charlot 2003). Now looking at the combined spectrum, the typical AGN features such as strong broad emission lines and blueshifted absorption lines are significant; in contrast, the starlight component is almost lost of any features, except for the appearance of weak Ca H & K and NaI D absorption lines. On the other hand, the high FIR luminosity betrays recent violent star formation activities (see Section 2.2), suggestive of the presence of a (obscured) young stellar population. Meanwhile, both the large extinction in the UV and the large Balmer decrement of narrow hydrogen emission lines (see Section 2.4.2 and Table 4) suggest the star formation region is enshrouded by thick dust. Here we can roughly estimate the lower limit to the extinction for the young stellar population that ionizes the HII region and powers the nebular emission lines and FIR dust emission. Using the SFR estimated from PAH (see Section 2.5) and the relation SFR($M_{\odot}$ yr$^{-1}$) = $7.9 \times 10^{-42}$L(Hα) (erg s$^{-1}$) (Kennicutt 1998), we get the predicted luminosity of Hα for the HII region, 1.78 x 10$^{43}$ erg s$^{-1}$. Yet the observed luminosity of the
narrow H$\alpha$ line $L$ (narrow H$\alpha$) is only $1.88 \times 10^{41}$ erg s$^{-1}$. Assuming all the emission of the narrow H$\alpha$ component is from star formation, then the extinction for narrow H$\alpha$ emission line is $A_{H\alpha} = 4.9$; applying the LMC extinction curve, the $E_{B-V}$ of the young stellar population is thus $\sim 2.6$. In addition, the excess emission between $\sim 2$ and $10 \mu m$, which

Figure 2. SDSS images of J1634+2049 and the 2D imaging decomposition by GALFIT. From top to bottom are SDSS $g$, $r$, and $i$ bands, and the composite of the three bands, respectively. The left column shows the original image, the middle column shows the GALFIT model (PSF + Sérsic), and the right column shows the residual image. All images are oriented with north up and east to the left; the black line marks a scale of $10^\prime$ ($\sim 23.1$ kpc). The green polygon denotes the ring region, which as well as the foreground stars has been masked out in GALFIT fitting.
appears in this spectrum, has been widely regarded to be originated from the hot dust of \( \sim 1500 \) K in the studies of AGN SEDs (Rieke 1978; Edelson & Malkan 1986; Barvainis 1987; Elvis et al. 1994; Glikman et al. 2006).

Based on the above analysis, we can decompose the continuum (3000–22000 Å in rest frame wavelength) with the following model:

\[
F_{\lambda} = C_{\text{nucleus}} A_{\text{nucleus}} (E_{\text{B–V}}^{\text{nucleus}}, \lambda) \lambda^\alpha + C_{\text{bb}} B_{\lambda} (T_{\text{dust}}) \\
+ C_{\text{host,1}} SSP (\geq 2 \text{ Gyr}) \\
+ C_{\text{host,2}} A_{\text{host}} (E_{\text{B–V}}^{\text{host}}, \lambda) SSP (\leq 1 \text{ Gyr}),
\]

where \( F_{\lambda} \) is the observed spectrum in the rest frame, \( C_{\text{nucleus}}, C_{\text{host,1}}, C_{\text{host,2}}, \) and \( C_{\text{bb}} \) are the factors for the respective components; \( A_{\text{nucleus}} \) and \( A_{\text{host}} \) are the dust extinction to the AGN emission and to the young stellar population, respectively; and \( B_{\lambda} (T_{\text{dust}}) \) is the Planck function. As \( \alpha \) (the intrinsic AGN continuum slope) and \( E_{\text{B–V}}^{\text{nucleus}} \) are somehow degenerate, in the fitting \( \alpha \) is fixed to \( \sim 1.7 \) while \( E_{\text{B–V}}^{\text{nucleus}} \) is free, which is the common recipe for reddened AGN continua in the literature (see, e.g., Dong et al. 2005, 2012, and Zhou et al. 2006). The two stellar populations are modeled by two simple stellar population (SSP) templates from Bruzual & Charlot (2003), with the metallicity being fixed to the solar (Z = 0.02) for simplicity. A SSP is defined as a stellar population whose star formation duration is short compared with the lifetime of its most massive stars. Based on the analysis in the preceding paragraph, we select 30 SSP templates with ages between 50 Myr and 1 Gyr to model the young stellar component and 28 SSP templates with ages between 5 and 12 Gyr to model the old stellar component. We traverse every possible SSP template in the library during the fitting, and use the IDL routine MPFIT (Markwardt 2009) to do the job with \( C_{\text{nucleus}}, C_{\text{host,1}}, C_{\text{host,2}}, C_{\text{bb}}, E_{\text{B–V}}^{\text{nucleus}}, E_{\text{B–V}}^{\text{host}}, \) and \( T_{\text{dust}} \) being free parameters. Detailedly, as the average QSO spectrum reddened with LMC extinction curve by \( E_{\text{B–V}} = 0.66 \) fits this object well (see Figure 1), the \( E_{\text{B–V}}^{\text{nucleus}} \) is initially set to be 0.66 and allowed to vary freely within 0.3–0.8 according to the Balmer decrement measured from the broad-line H\( \alpha / H\beta \). The \( E_{\text{B–V}}^{\text{host}} \) is initially set to be 2.6, varying within 1.0–3.0 according to the Balmer decrement.
measured from the narrow-line H\textalpha/H\beta (see Table 4). We also set the fraction of the power-law component in the r band to be 0.5 initially and allow the value varying within 0.4–0.6.

The fitting converges on the two SSP templates with ages of 127 Myr and 9 Gyr, respectively, \(E_{B-V}^{\text{nucleus}} = 0.41, E_{B-V}^{\text{host}} = 2.2, \) and \(T_{\text{dust}} = 1394\) K that is close to the result of Glikman et al. (2006). As Figure 4 shows, the fractions of the decomposed nucleus and starlight components in the spectrum are basically consistent with the imaging decomposition. The blackbody continuum from the hot dust of the presumable torus should be the dominant emission in the K band and the WISE W1 band. To check it, we extend the fitting result to the WISE W1 band and find that it can well reproduce the observed data point. As to the best-fit starlight spectrum composed of the two SSP components, it reproduces well the Ca II H & K and H\textalpha absorption lines, yet it overestimates the Na I D absorption.
This discrepancy may be due to the uncertainty of the decomposition or more likely the contamination to the Na D absorption by the nearby He I λ5876 emission line. The concern on the decomposition is mainly about the decomposition of the two stellar components. To assess the reliability of the decomposed stellar components (mainly of their ages), we do the following checks. First, regardless of the above physical arguments to justify the use of two SSPs, we test this point with the data (lest the spectral quality should not be sufficient to support it). Using a single SSP with the metallicity and age being free parameters for the starlight in the model, we obtain the best fit with a much larger minimum \( \chi^2 \). The reduced \( \chi^2 \) increases by \( \Delta \chi^2 = 2.5 \) and the two-SSP model is favored according to F-test (the chance probability \( P_{null} = 0 \)). Then we check the algorithm of MPFIT for global minimum, as follows. (1) We loop over the grid of the 28 templates for the old stellar population component and for every case of the assigned template for the old component we get the best-fit young stellar component according to minimum \( \chi^2 \); these 28 best-fit young populations have ages in the range of 80–210 Myr. (2) On the other hand, we loop over the grid of the 30 templates for the young stellar population component and for every case we get the best-fit old component by \( \chi^2 \) minimization; these 30 best-fit old stellar populations have ages in the range of 8–11 Gyr. We can see that, at least, the ages of the two components can be well separated. (3) Furthermore, considering that the Ca II H & K and He I absorption lines are dominated by the old stellar component, we devise instead a \( \chi^2_{abs} \) as calculated in the spectral region of 3900–4050 Å to better constrain the fitting of the old stellar component. We repeat the procedure of (2) yet with minimizing \( \chi^2_{abs} \). For every case of the held template for the young component, the best-fit SSP template for the old component may be different from (2). Yet the ages of the best-fit old components of all the 30 cases are still in the range of 8–11 Gyr. Certainly, we should note that the above checks do not account for the effect of the internal dust extinction parameters (see Equation (1)), which should impact the fitting of the two stellar components.

Since for any SSP the age and metallicity are degenerate parameters, we also try to model the starlight with two SSPs with their age and metallicity set to be free (allowing \( Z = 0.0001, 0.0004, 0.0004, 0.0008, 0.02 \) and 0.05). This test scheme yields that the best-fit two SSPs have ages of 47.5 Myr and 12 Gyr, and both converge to an extremely low metallicity \( Z = 0.0004 \). Such a metallicity is even lower than that of most of the most metal-poor (dwarf) galaxies, which is unrealistic for J1634+2049 with a mass/size similar to the Milky Way. Besides, this scheme does not change much the best-fit starlight (the sum of the two stellar populations): the difference of the starlight between this scheme and the solar metallicity scheme yields that the best-fit young component according to minimum \( \chi^2 \) essentially exhibit a narrow peaked cusp, the two components are modeled with a single Gaussian each, and the broad component is fitted with a Gaussian each; more additional Gaussian(s) can be added into the model for a component if the \( \chi^2 \) decreases significantly with an F-test probability \( \leq 0.05 \). A first-order polynomial is adopted to fine-tune the local continuum. The fitting turns out, with a reduced \( \chi^2 \) of 0.99, that the narrow component and the blueshifted cusp are well fitted with a single Gaussian each, and the broad component is sufficiently fitted with two Gaussians (see Figure 6).

The Pγ emission line is heavily blended with the He I λ10830 emission line. In addition, two broad He I λ10830 absorption troughs are located in the blue wing of the He I λ10830 emission line, which increases the complexity of decomposing this blend. We use the Po as the template to fit Pγ. To be specific, we assume that Pγ has also three components, namely the narrow, broad, and blueshifted cuspy ones, and every component shares the same profile as the corresponding one of Po, only with free intensity factors.\(^\text{13}\) The profile of the He I λ10830 emission line is different from that of Po, except that its blueshifted cusp is located at a similar blueshifted velocity to the cusps of Po and Hβ. We understand that He I λ10830 is a high-ionization line (with an ionization potential of 24.6 eV), and it is well known that high-ionization lines generally have a more complex profile than low-ionization lines (e.g., C IV λ1549, Wang et al. 2011; see also

\(^\text{13}\) Except the width of the narrow Pγ, for which we adopt the fitting result with it set to be free (cf. Table 3). This is because the narrow Pγ component stands high over the broad component and the fitting can be significantly improved by relaxing its width from being tied to that of narrow Po.

2.4.2. Emission Lines

As the Figure 4 shows, strong emission lines of J1634+2049 display intensively in four spectral regions: Hβ + [O III] λλ4959, 5007 + Fe II multiplets, Hα + [N II] λλ6548, 6583 + [S II] λλ6716,6731, Pγ + He I λ10830, and Po (note that the Pβ and Pθ emission lines are of low S/N).

Before we fit the four regions, we first take a look at the profiles of the various emission lines. Figure 5 shows Po, Pγ, He I λ10830, Hα, He I λλ5876, [O III] λλ5007, Hβ, Ne II λ3869, and [O III] λλ3727 emission lines in the common velocity space. We note the following major points. (1) The total profiles of the recombination lines (such as Po, Pγ, He I λ10830, Hα, He I λλ5876, and Hβ) apparently exhibit a narrow peaked component and a lower and broad base; the two components appears separable from each other. (2) Besides, probably of the most interesting, in Hβ and He I λλ5876, 10830, there clearly exists an extra cuspy narrow component that is blueshifted by a velocity of \( \sim 900 \) km s\(^{-1}\); the blueshift velocity is consistent in the three lines. (3) The whole profile of high-ionization narrow forbidden emission lines, [O III] λλ5007 and Ne II λ3869, is blueshifted evidently, with an offset velocity of \( \approx 500 \) km s\(^{-1}\) according to the peak of [O III] λλ5007; this kind of bulk blueshifting seems to be present in the low-ionization forbidden line [O III] λλ3727, yet with a smaller blueshift.

We start the line profile fittings with the Po emission line, which is the strongest of the hydrogen Paschen series, and basically free of blending (unlike He I λ10830). Its profile can be used as a template to fit the Pγ + He I λ10830 blends (Landt et al. 2008). The high contrast between the narrow peak and broad base of the Po profile makes it able to decompose easily into the broad and narrow components. There is another weak yet significant excess cusp on the blue side of the Po profile. If this cusp belongs to Po, its velocity offset is almost the same as the narrow cusp of Hβ mentioned above (at velocity of \( \approx 900 \) km s\(^{-1}\) in Figure 5). Therefore, we take this cusp as the third component of the Po profile in the model. Every components are modeled with (multiple) Gaussians. Initially, the narrow, the blueshifted cuspy, and the broad components are fitted with one Gaussian each; more additional Gaussian(s) can be added into the model for a component if the \( \chi^2 \) decreases significantly with an F-test probability \( \leq 0.05 \). A first-order polynomial is adopted to fine-tune the local continuum. The fitting turns out, with a reduced \( \chi^2 \) of 0.99, that the narrow component and the blueshifted cusp are well fitted with a single Gaussian each, and the broad component is sufficiently fitted with two Gaussians (see Figure 6).
This is interpreted by the presence of pronounced other components in high-ionization lines—particularly the emission originated from the AGN outflows—in addition to the normal component emitting from the virialized BLR clouds that is located at the systematic redshift (see, e.g., Zhang et al. 2011). To get a convergent fitting for the He I line, we use the profile of the broad Pα (namely the two-Gaussian model) as the template to model the virialized broad He I component, a Gaussian to model the NLR-emitted He I, a Gaussian to the blueshifted cusp, and as many more Gaussians as statistically guaranteed (namely F-test probability ≤ 0.05) to account for the remaining flux. In the fitting the absorption...
The best-fit model is shown in the lower left panel of Figure 6 with a reduced $\chi^2 = 1.15$; the best-fit parameters as listed in Table 3. Besides the virialized broad component, the NLR one and the cusp, finally there are two additional Gaussians to account for an extra blueshifted broad component. This extra component is blueshifted by approximately $1360$ km s$^{-1}$, which is consistent with the aforementioned outflow interpretation for the profiles of high-ionization broad lines. Note that such a blueshifting is not a direct identification but merely ascribed to the asymmetry of the broad-line He I$\lambda10830$ profile, which is different from the situation of the blueshifted narrow cusp component. Hereafter when necessary, for the ease of narration we denote this blueshifted broad component with “outflowB”, and the blueshifted narrow cusp “outflowN.”

In the optical, H$\alpha$ shows a strong broad base and an apparent narrow peak, which are blended with [N II]$\lambda\lambda6548, 6583$ doublet. The red wing of the broad base is also slightly affected by [S II]$\lambda\lambda6716, 6731$ doublet. H$\beta$, as we stress in the above, shows an apparent narrow cusp blueshifted by approximately $900$ km s$^{-1}$, which is consistent with the extra cuspy components revealed in P$\alpha$ and He I$\lambda10830$ (see Figure 5). This suggests that H$\alpha$ should also have such a blueshifted cuspy component. Because the Balmer lines are heavily blended with strong Fe II multiplet emission, we fit the continuum-subtracted spectrum (namely, simultaneously fitting Balmer lines $+[O III] + [N II] + [S II] +$ Fe II), following the methodology of Dong et al. (2008). Specifically, we assume the broad, narrow, and blueshifted cuspy components of H$\beta$ and H$\gamma$ have the same profiles as the respective components of H$\alpha$. The $[O III]$$\lambda\lambda4959, 5007$ doublet lines are assumed to have the same profile and fixed in separation by their laboratory wavelengths; the same is applied to [N II]$\lambda\lambda6583, 6548$ doublet lines and to [S II]$\lambda\lambda6716, 6731$ doublet lines. The flux ratio of the $[O III]$ doublet $5007/4959$ is fixed to the theoretical value of 2.98 (e.g., Storey & Zeippen 2000; Dimitrijević et al. 2007); the flux ratio of the [N II] doublet $6583/6548$ is fixed to the theoretical value of 2.96 (e.g., Acker et al. 1989; Storey & Zeippen 2000). We use Gaussians to model every component of the above emission lines as we describe in the above for the NIR lines, starting with

![Figure 6](image-url)
Table 3

| Emission Line | Centroid<sup>a</sup> (Å) | FWHM<sup>b</sup> (km s<sup>-1</sup>) | Flux<sup>c</sup> (10<sup>-17</sup> erg s<sup>-1</sup> cm<sup>-2</sup>) |
|---------------|------------------------|----------------|----------------------------|
| [O iii]λ3727  | 3727.39 ± 0.21         | 491 ± 5        | 184 ± 3                     |
| (narrow)      |                        |                |                            |
| [O iii]λ3727  | 3720.19 ± 2.22         | 738 ± 5        | 89 ± 8                     |
| (outflow)     |                        |                |                            |
| Hγ(broad)<sup>d</sup> | 4342.43                | 2955           | 487 ± 25                   |
| Hγ(narrow)<sup>d</sup> | 4340.74                | 343            | 8 ± 5                      |
| Hγ(outflow)<sup>d</sup> | 4327.58                | 737            | 7 ± 9                      |
| Hβ(broad)<sup>d</sup> | 4683.52                | 2955           | 1378 ± 14                  |
| Hβ(narrow)<sup>d</sup> | 4861.63                | 343            | 42 ± 3                     |
| Hβ(outflow)<sup>d</sup> | 4846.50                | 737 ± 20       | 58 ± 5                     |
| [O ii]λ5007  | 4992.80 ± 0.32         | 1707 ± 18      | 1043 ± 12                  |
| (narrow)<sup>e</sup> |                        |                |                            |
| Hα           | 6299.87 ± 0.53         | 421            | 61 ± 6                     |
| HeⅡ(broad)   | 6565.74 ± 0.18         | 2955 ± 12      | 8719 ± 66                  |
| HeⅡ(narrow)  | 6563.18 ± 0.03         | 343 ± 5        | 427 ± 8                    |
| HeⅡ(outflow) | 6543.30 ± 0.13         | 737            | 351 ± 17                   |
| [N ii]λ6583  | 6583.05 ± 0.03         | 370 ± 4        | 458 ± 6                    |
| [S ii]λ6716  | 6716.42 ± 0.12         | 421 ± 10       | 148 ± 5                    |
| (narrow)<sup>f</sup> |                        |                |                            |
| HeⅠλ10830   | 10840.33 ± 1.48        | 3803           | 994 ± 53                   |
| (broad)<sup>g</sup> |                        |                |                            |
| HeⅠλ10830   | 10831.33 ± 0.45        | 178 ± 12       | 162 ± 25                   |
| (narrow)<sup>h</sup> |                        |                |                            |
| HeⅠλ10830   | 10801.14 ± 0.33        | 140 ± 23       | 21 ± 7                     |
| (outflow)<sup>i</sup> |                        |                |                            |
| HeⅠλ10830   | 10782.34 ± 0.33        | 1612 ± 23      | 764 ± 37                   |
| (outflow)<sup>j</sup> |                        |                |                            |
| P(1/2)<sup>l</sup> | 10948.39               | 3803           | 874 ± 34                   |
| P(1/2)(narrow)<sup>l</sup> | 10904.50 ± 0.31       | 204 ± 18       | 69 ± 10                    |
| P(1/2)(outflow)<sup>l</sup> | 10907.00              | 575            | 50 ± 11                    |
| Pα(broad)<sup>l</sup> | 18768.64 ± 0.20        | 3803 ± 103     | 2484 ± 79                  |
| Pα(narrow)<sup>l</sup> | 18755.96 ± 0.19        | 310 ± 11       | 486 ± 13                   |
| Pα(outflow)<sup>l</sup> | 18697.69 ± 1.64        | 575 ± 75       | 146 ± 16                   |

Notes.

- <sup>a</sup> Vacuum rest frame wavelength.
- <sup>b</sup> Corrected for instrumental broadening.
- <sup>c</sup> Adopting the profile of the broad component of Hα.
- <sup>d</sup> Adopting the profile of the narrow component of Hα.
- <sup>e</sup> Adopting the profile of the outflow component of Hα.
- <sup>f</sup> Adopting the profile of the broad component of Hβ.
- <sup>g</sup> Adopting the profile of the narrow component of Hβ.
- <sup>h</sup> Adopting the profile of the outflow component of Hβ.
- <sup>i</sup> Adopting the profile of [S ii]λ6716.
- <sup>j</sup> Adopting the profile of the broad component of Pα.
- <sup>k</sup> Adopting the profile of the outflow component of Pα.

One Gaussian and adding in more if the fit can be improved significantly according to the F-test. The best-fit model turns out that two Gaussians are used for the broad component of the Hydrogen Balmer lines; one for the narrow component and one for the blueshifted cusp. Two Gaussians are used for every line of the [O iii] doublet and one for all the other aforementioned narrow lines. The Fe ii multiplet emission is a model with the two separate sets of analytic templates of Dong et al. (2008), one for the broad-line Fe ii system and the other for the narrow-line system, constructed from measurements of I Zw 1 by Véron-Cetty et al. (2004). Within each system, the respective set of Fe ii lines is assumed to have no relative velocity shifts and the same relative strengths as those in I Zw 1. We assume that the broad and narrow Fe ii lines have the same profiles as the broad and narrow Hβ, respectively; see Dong et al. (2008, 2011) for the detail and justification. A first-order polynomial is adopted to fine-tune the local continuum of the Hα + [N ii] + [S ii] region and the Hβ + Hγ region, respectively. The best-fit model is presented in Figure 6, with a reduced χ² of 2.34 in the Hβ + Hγ region and a reduced χ² of 1.46 in Hα + [N ii] + [S ii] region. The somehow large reduced χ² in the Hβ + Hγ region is mainly due to the excess emission in the red wing of Hβ, the so-called “red-shelf” commonly seen in type-1 AGNs and has been discussed in the literature (e.g., Meyers & Peterson 1985; Véron et al. 2002). It is probably the residual of Fe ii multiplet 42 (λ4924, 5018, 5169), or broad He iλ4922,5016 lines (see Véron et al. 2002), or just the mis-match between Hα and Hβ. Since it is irrelevant to the components of interest in this work, we do not discuss it further. We also fit [O ii]λ3727 line, which is well isolated and easily fitted with two Gaussians (Figure 5).

The measured line parameters are listed in Table 3. The extinction of the broad, narrow, and blueshifted cuspy components of the Balmer lines can be derived from the observed Balmer decrement Hα/Hβ. The intrinsic value of the broad-line Hα/Hβ is 3.06 with a standard deviation of 0.03 dex Dong et al. (2008). A value of 3.1 is generally adopted for the intrinsic narrow-line Hα/Hβ in AGN (Halpern & Steiner 1983; Gaskell & Ferland 1984). The intrinsic Pα/Hβ of AGNs is very close to the Case-B value 0.34 (Gaskell & Ferland 1984).

In Table 4 we list the observed Hα/Hβ and Pα/Hβ as well as E<sub>b-v</sub> assuming the extinction curve of the LMC (R<sub>b</sub> = 2.6). The extinctions to the broad components and to the blueshifted cusps are similar, while the narrow components suffer much larger extinction, indicating that the NLR could be more dust-obsured.

Table 4

| Decrement | Component | Value<sup>l</sup> |
|-----------|-----------|------------------|
| Hα/Hβ<sup>d</sup> | broad | 6.33 ± 0.08 |
| Hα/Hβ<sup>d</sup> | narrow | 10.17 ± 0.75 |
| Hα/Hβ<sup>d</sup> | outflow (narrow) | 6.05 ± 0.60 |
| Pα/Hβ<sup>d</sup> | broad | 1.80 ± 0.06 |
| Pα/Hβ<sup>d</sup> | narrow | 11.57 ± 0.90 |
| Pα/Hβ<sup>d</sup> | outflow (narrow) | 2.52 ± 0.35 |

Notes.

- <sup>a</sup> Intrinsic Hα/Hβ for BLR is 3.06 (Dong et al. 2008).
- <sup>b</sup> Intrinsic Hα/Hβ for NLR and outflowing gas is 3.1 (Dong et al. 2008).
- <sup>c</sup> Intrinsic Pα/Hβ is 0.34 (Gaskell & Ferland 1984).

Notes.

2.4.3 Absorption Lines

J1634+2049 shows He iλ3889, 10830 and Na i D BALs, which are generally deemed to be caused by AGN outflows. To analyze the absorbed intensities, we need to first identify the pre-absorption AGN spectrum and normalize the data with it. We first subtract the best-fit narrow emission lines and the starlight component from the observed spectrum before normalization, as the absorption gas does not cover the NLR in all well studied BALs. There are three components left in the observed spectrum: the power-law continuum, the broad emission lines, and the blackbody (hot dust) continuum. He i absorption might be normalized in different ways depending on whether the absorbing gas covers the torus and/or the BLR. We notice that the He iλ10830 absorption trough shows a flat bottom, which indicates that the He iλ10830 line is saturated and the residual fluxes should
be zero in those pixels. Meanwhile, we find the residual fluxes at the line centroid of He I $\lambda 10830$ is close to zero after subtracting the starlight continuum (see the middle panel in Figure 4). Again, we check carefully the decomposition of the continuum. The fractions of the power law and starlight in the optical are determined by the decomposition of images, which should be reliable. The fitted blackbody continuum, which originates from the hot dust of the torus, is the dominant emission in the $K$ band and the WISE W1 band, which also accords with the expectation. Therefore, we conclude that the absorbing gas is likely exterior to the torus.

The left panels of Figure 7 demonstrate the absorption profiles of He I $\lambda\lambda 10830, 3889$ and Na I D absorption lines in velocity space. On the SDSS spectrum, the He I $\lambda 3889$ absorption is clearly detected. It splits into two absorption troughs, including a larger one near $-4100$ km s$^{-1}$ and a second one near $-3400$ km s$^{-1}$, totally covering the velocity ranges $-5000$ to $2800$ km s$^{-1}$ (trough A and B in Figure 7). The strong He I $\lambda 10830$ absorption line is identified on the TripleSpec spectrum. It separates into two major troughs spreading from $v \sim -5000$ km s$^{-1}$ to $v \sim -2000$ km s$^{-1}$ (trough A+B, C in Figure 7). He I $\lambda 3889$ and He I $\lambda 10830$ are transitions from the same lower level, so they should have the same velocity profile theoretically. The strongest trough, which corresponds to the trough A and B of He I $\lambda 3889$, is obvious saturated, since the bottom of the absorption line is flat. This can be naturally explained by the large optical depths ratio of He I $\lambda 10830, 3889$.

We use the Voigt profile (Hjerting 1938; Carswell et al. 1984) to fit these absorption troughs, which are shown in Figure 7. The Voigt profile is implemented with the program $x\_voigt$ in the XIDL package. He I $\lambda\lambda 3889, 10830$ are the strongest two transitions from the metastable state to the 2p, 3p states, and the $\tau$ ratio ($\propto f_{\lambda} N_{\text{ion}}$) of He I $\lambda\lambda 10830, 3889$ is

\footnote{http://www.ucolick.org/~xavier/IDL/}
23.5:1 (Liu et al. 2015, their Table 2). If the lines are not saturated and the absorbers fully cover the source, the normalized flux \( R_{10830} = R^{23.5}_{3889} \). The cyan dotted lines in the upper left panel of Figure 7 shows the \( \text{He} \, \lambda \lambda 10830,3889 \) absorption profile predicted from the \( \text{He} \, \lambda \lambda 10830,3889 \) absorption trough under the full-coverage assumption. The red wing of the observed \( \text{He} \, \lambda \lambda 10830 \) absorption line fits the predicted profile well, while the blue wing of the observed profile is evidently different from the predicted. This suggests a full-coverage situation for the absorbing gas of low outflow speed and a partial coverage for the gas of high speed.

Based on the above inference, we try to estimate the covering factor of the outflow gas assuming a simple partial-coverage model, where the observed normalized spectrum can be expressed as follows:

\[
R = \left(1 - C_f(v)\right) + C_f(v)e^{-\tau(v)}. \tag{2}
\]

Here the \( \tau \) ratio of \( \text{He} \, \lambda \lambda 10830,3889 \) is 23.5. Although the spectral resolutions of the SDSS and TripleSpec spectrum are not high enough for us to study the velocity structure of absorption trough in detail, the tendency and the mean of \( C_f \) are reliable. We bin the spectrum by three pixels and perform the calculation and analysis using the binned data. Following the methodology of Leighly et al. (2011), we derive the covering fraction \( (C_f) \), the optical depth \( (\tau) \), and the column density of \( \text{He} \, \lambda \), as a function of velocity (see right panels of Figure 7), the covering factor of outflow with \( v < 4300 \text{ km s}^{-1} \) is close to 1.0, contrasting with the high speed outflow with \( v > 4300 \text{ km s}^{-1} \). The average covering factor of component A is 0.82, and \( \log N_{\text{He} \, \lambda} \) \( (\text{cm}^{-2}) = 15.01 \pm 0.16 \). Assuming \( \text{Na} \, \text{D} \) absorption lines have the same covering factor with \( \text{He} \, \lambda \) absorption lines, then we get the ionic column density of \( \text{Na} \, \text{I} \) is \( \log N_{\text{Na} \, \lambda} \) \( (\text{cm}^{-2}) = 13.31 \pm 0.21 \).

2.5. Analysis of the Spitzer Spectrum

The MIR spectrum of J1634+2049 shows significant PAH emission features and a steeply rising continuum toward long wavelength end (see Figure 8). We measured the apparent strength (namely the apparent optical depth) of the 9.7 \( \mu \text{m} \) silicate feature following the definition by Spoon et al. (2007),

\[
S_{\text{all}} = \ln \frac{f_{\text{obs}(9.7 \mu \text{m})}}{f_{\text{cont}(9.7 \mu \text{m})}}, \tag{3}
\]

where \( f_{\text{obs}(9.7 \mu \text{m})} \) is the observed flux density at 9.7 \( \mu \text{m} \) and \( f_{\text{cont}(9.7 \mu \text{m})} \) is the continuum flux density at 9.7 \( \mu \text{m} \). Following Spoon et al. (2007), we estimate \( f_{\text{cont}(9.7 \mu \text{m})} \) from a power-law interpolation over 5.5–14.0 \( \mu \text{m} \). It gives \( S_{\text{all}} = 0.01 \pm 0.03 \) (1\( \sigma \)), indicating (almost) neither silicate emission nor absorption. It is interesting that the silicate absorption is not present at all, as the analysis in the UV and optical bands demonstrates instead that J1634+2049 is dust-obscured. In the diagnostic diagram of EW(PAH 6.2 \( \mu \text{m} \)) versus 9.7 \( \mu \text{m} \) silicate strength as shown in Figure 1 of Spoon et al. (2007), galaxies are located mainly around two branches closely: a diagonal and a horizontal. J1634+2049 belongs to the horizontal branch in the 1A class (EW(PAH 6.2 \( \mu \text{m} \)) = 0.050 \pm 0.003, see Table 5).
Interestingly, Mrk 231 also belongs to this class but with a larger 9.7 μm silicate absorption strength (S_{sil} ∼ 0.65) as well as a smaller PAH 6.2 μm EW (∼0.01). As discussed by Spoon et al. (2007), the two distinct branches reflect likely the differences in the spatial distribution of the nuclear dust. Galaxies in the horizontal branch may have clumpy dust distributions, which produce only shallow absorption features. In J1634+2049 as we analyzed above, the narrow emission lines suffer larger extinction than the broad emission lines and the young stellar component is obscured more seriously than the AGN continuum and broad/narrow emission lines. This implies that the dust in the nuclear region is clumpy and patchy; i.e., the nuclear region is not fully enshrouded by dust. Meanwhile, there are significant Ne II 12.8 μm and Ne II 15.6 μm emission lines, but no higher-ionization lines such as Ne v 14.3 μm and [O vi] 25.89 μm in the MIR spectrum; this may indicates that the AGN contributes insignificantly to the MIR emission (Farrah et al. 2007).

We use the PAHFIT spectral decomposition code (v1.2)\(^\text{15}\) (Smith et al. 2007) to fit the spectrum as a sum of dust attenuated starlight continuum, thermal dust continuum, PAH features, and emission lines. As the above analysis suggests that the AGN contribution in the MIR flux is low, we adopt the fully mixed extinction geometry. Figure 8 shows the MIR spectrum and the best-fit decomposition with a reduced χ^2 = 1.16. The τ_{60} is fitted to be ∼0. The fluxes of the main MIR components derived form the PAHFIT are listed in Table 5. The flux errors are given by PAHFIT; see Smith et al. (2007) for the detail. The errors of the EWs are estimated according to error propagation formula, where the uncertainties of continuum is estimated as follows. For the spectral region of every emission line feature, we calculate the residual between the raw spectrum and the best-fit (continuum + line) and then the standard deviation of the residual in this region is taken as the 1-σ error of the continuum placement. The total PAH Luminosity \( \log L_{\text{PAH}} \) (erg s\(^{-1}\)) = 43.64 ± 0.02, which is ∼1.2% of total IR luminosity of this source. We estimate SFR from the PAH features using the relation of Farrah et al. (2007), \( \text{SFR}(M_{\odot} \text{ yr}^{-1}) = 1.18 \times 10^{-44} L_{\text{PAH}} + 11.2 \mu m \), and the errors of the SFR derived using this equation are of order ∼50% for individual objects. According to our measurements, we get the SFR = 141 ± 71 M_{\odot} yr\(^{-1}\). As a double check, we also estimate the SFR from the 24 μm emission. Adapting the relation for galaxies with \( L_{\text{IR}} \geq 10^{11}L_{\odot} \) of Rieke et al. (2009), the error of the SFR derived from which is whitin 0.2 dex, we get SFR = 143.7 ± 53 M_{\odot} yr\(^{-1}\). The SFR values estimated from the PAH, 24 μm emissions and the total IR luminosity are well consistent, and they are also consistent with the upper limit of the SFR estimated from the IR luminosity. Hence, we adopt the SFR estimated from PAH in this paper.

3. RESULTS

3.1. Central Black Hole

With the measured luminosity and line width of the broad emission lines, we can estimate the mass of the central BH using the commonly used virial mass estimators. We use the broad Hα-based mass formalism given by Xiao et al. (2011, their Equation (6)), which is based on Greene & Ho (2005b, 2007) but incorporates the recently updated relation between BLR size and AGN luminosity calibrated by Bentz et al. (2009). The broad Hα luminosity is corrected for the broad-line extinction using the LMC extinction curve, resulting \( L_{\text{H}_\alpha} = 1.5 \times 10^{43} \text{erg s}^{-1} \). Together with FWHM(\text{H}_\alpha) = 2955 km s\(^{-1}\), the central black hole mass is estimated to be \( M_{\text{BH}} = 7.94 \times 10^7 M_{\odot} \).

We calculate monochromatic continuum luminosity \( L_\lambda(5100 \text{ Å}) \) at 5100 Å from the best-fit power-law component (see Section 2.4.1). The best-fit \( E_{B-V} \) value of the power-law component is 0.42 (assuming the LMC extinction curve) and the extinction corrected luminosity \( L_\lambda(5100 \text{ Å}) = 2.47 \times 10^{44} \text{erg s}^{-1} \). As a check, we also estimate the \( L_\lambda(5100 \text{ Å}) \) from the Hα flux (Greene & Ho 2005b), which gives \( L_\lambda(5100 \text{ Å}) = 2.49 \times 10^{44} \text{erg s}^{-1} \), fairly consistent with the above one obtained from the best-fit power law. Then we calculate the bolometric luminosity using the conversion \( L_{\text{bol}} = 0.75 \times 4.89 + 0.91 L_\lambda(5100 \text{ Å}) \) (Runnoe et al. 2012), which gives \( L_{\text{bol}} = 1.45 \times 10^{45} \text{erg s}^{-1} \). The corresponding Eddington ratio is thus \( L_{\text{bol}}/L_{\text{Edd}} = 0.15 \). Based on the bolometric luminosity, the amount of mass being accreted is estimated as follows:

\[
M_{\text{sec}} = \frac{L_{\text{bol}}}{\eta c^2} = 0.26 M_{\odot} \text{yr}^{-1},
\]

where we assumed an accretion efficiency \( \eta \) of 0.1, and \( c \) is the speed of light.

3.2. Host Galaxy

The 2D decomposition of the SDSS images (Section 2.3) yields that the host galaxy is an early-type galaxy with Sérsic index \( n = 4 \) and \( M_V = -22.07 \text{ mag} \). We use the K-band luminosity of the starlight to calculate the stellar mass of J1634+2049, which is relatively insensitive to dust absorption and to stellar population age. Into & Portinari (2013) provide a calibration of mass-to-light ratios against galaxy colors, and we here adopt their Table 3 relation for log \( M/L_{Ks} = 1.055 (B-V) - 1.066 \) with a scatter of ±0.13 dex. The \( B, V \) magnitude and \( L_{Ks} \) of J1634+2049 is calculated by convolving

\[
\text{Table 5}
\]

| Emission Line | Intensity* (10\(^{-20}\) W cm\(^{-2}\)) | EW* (μm) |
|--------------|-----------------|---------|
| PAH 6.2 μm | 1.47 ± 0.09 | 0.050 ± 0.003 |
| PAH 7.7 μm complex | 3.06 ± 0.16 | 0.135 ± 0.007 |
| PAH 8.3 μm | 0.57 ± 0.12 | 0.028 ± 0.006 |
| PAH 8.6 μm | 1.08 ± 0.08 | 0.055 ± 0.004 |
| PAH 11.3 μm complex | 1.25 ± 0.08 | 0.085 ± 0.005 |
| PAH 12.0 μm | 0.10 ± 0.09 | 0.007 ± 0.006 |
| PAH 12.6 μm complex | 1.44 ± 0.14 | 0.106 ± 0.010 |
| PAH 13.6 μm | 0.43 ± 0.10 | 0.033 ± 0.008 |
| PAH 16.4 μm | 0.13 ± 0.07 | 0.011 ± 0.006 |
| PAH 17 μm complex | 0.34 ± 0.25 | 0.029 ± 0.019 |
| [Ne ii]λ12.8 μm | 0.32 ± 0.03 | 0.034 ± 0.004 |
| [Ne ii]λ15.6 μm | 0.12 ± 0.03 | 0.014 ± 0.004 |

Notes.

* Uncertainties are given by the PAHFIT code; see Smith et al. (2007) for the detail.

* Uncertainties are estimated according to error propagation formula.

\(^{15}\) http://tir.astro.utoledo.edu/jdsmith/research/pahfit.php
the decomposed stellar component (see 2.4.1) with the Jonson B band and V band response curve. We get B – V = 0.63 and 
$L_K = 2.4 \times 10^{42}$ erg s$^{-1}$, the stellar mass estimated from $L_K$ is $M_{\text{host}} = (1.8 \pm 0.54) \times 10^{11} M_\odot$. The error estimate accounts for the uncertainty of decomposition of starlight in K band (see Section 2.4.1) and the uncertainty of $M/L_K$. As a check, we also use V band luminosity to estimate the stellar mass. The UV–NIR continuum decomposition (Section 2.4.1) shows that the galaxy is dominated by an old stellar population with an age of $\approx$9 Gyr, which corresponds to a mass-to-light ratio of $M/L_V \approx 5.2 \ M_\odot/L_{V,\odot}$ (Bruzual & Charlot 2003). The V-band luminosity of the old stellar population is calculated by convolving the decomposed old stellar component with Jonson V-band response curve, which gives $L_V = 9.24 \times 10^{43}$ erg s$^{-1}$. Thus, the stellar mass is $M_{\text{host}} \approx 1.34 \times 10^{11} M_\odot$, which is basically consistent with stellar mass estimated from K-band luminosity.

Both the decomposition of the UV to NIR spectrum and the analysis of its SED and MIR spectrum reveal that there is a heavily obscured young stellar component (see Section 2.4.1) relating to the recent violent (obscured) star formation activities. This stellar component is heavily obscured in the UV and optical bands. We cannot see any sign of star formation activities from the optical image either. According to the decomposition of the optical continuum, this component accounts for $\leq$5% of the total continuum emission at 5100 Å. The best-fit result of the spectral decomposition gives the extinction of this component is $E_{B-V} = 2.2$ (see Section 2.4.1). That is why we can only reliably infer the presence of this young stellar component from the PAH emission in the Spitzer spectrum and the high FIR luminosity.

It is interesting to note the fact that the AGN narrow emission lines suffer more dust obscuration than the AGN broad lines, and the young stellar component suffer more than the two. This may tell us some clues of the spatial distribution of the dust. The AGN NLR is probably related to the heavily obscured H\textsc{ii} region. Next we investigate the situations of the dust. The AGN NLR is probably related to the heavily obscured H\textsc{ii} region. As we cannot measure the stellar velocity dispersion from the optical spectra, we use instead the velocity dispersion of the line-emitting gas in the NLR as traced by the low-ionization [S\textsc{ii}]$\lambda\lambda$6716, 6731 lines as a surrogate (Greene & Ho 2005a; Komossa & Xu 2007). For $\sigma_\alpha \approx \sigma_{\text{H}\beta} = 178.8$ km s$^{-1}$, the $M_{\text{BH}}-\sigma_\alpha$ relation for early-type galaxies of McConnell & Ma (2013, their Table 2) predicts $M_{\text{BH}} = 1.37 \times 10^8 M_\odot$, which is 1.6 times the virial BH mass based on the broad H\textsc{ii} line. Now we take the Sérsic $n = 4$ component (Section 2.3) as the galactic bulge, the $M_{\text{BH}}$–$L_{\text{bulge}}$ relation for early-type galaxies of (2013, their Table 2) predicts $M_{\text{BH}} = 3.47 \times 10^8 M_\odot$, which is ~4.4 times the virial estimate based on the broad H\textsc{ii} line. Considering the large uncertainties associated with these quantities (both methodologically and statistically), at this point we can only say that both the $M_{\text{BH}}-\sigma_\alpha$ and $M_{\text{BH}}$–$L_{\text{bulge}}$ relationships of the nearby inactive galaxies seem to hold in this object.

3.3. The Companion Galaxy and the Galactic Ring

There are two small faint galaxies to the west of the ring (C1 and C2 in Figure 9). The projected distances from the center of J1634+2049 to the centers of C1 and C2 are 18.5’ and 19.1’, corresponding to 35.7 kpc and 44.0 kpc, respectively. We spectroscopically observed the C1 and C2 galaxies using YFOSC mounted on Lijiang GMG 2.4 m telescope. The right panel of Figure 9 shows the obtained spectra of J1634 +2049 and C1. The C2 galaxy is too faint to get an effective spectrum. In the right top panel of Figure 9, the spectra of the Lijiang 2.4 m telescope and of the SDSS are compared to make sure that the wavelength calibration is reliable, which is key to determine the redshift of C1. The spectrum of C1 galaxy shows characteristics of an early-type galaxy (ETG) of old stellar population, with visible stellar absorption features and no emission lines. By comparing the C1 spectrum with the template spectra in the SWIRE template library (Polletta et al. 2007), a best-matched template of elliptical galaxy of 5 Gyr old is picked up. Because of no strong emission lines, we perform a grid search of redshift for the redshift of C1. The searching procedure is as follows. The redshift grids are set to be $0.11 < z < 0.15$. In every redshift grid, the template spectrum is brought to the observer frame, and is reddened with the Milky Way extinction along the direction of C1; then we fit the observed C1 spectrum with the reddened template and calculate the $\chi^2$. In this way we obtain the curve of $\chi^2$ with redshift, as shown in the bottom panel of Figure 9. The redshift value corresponding to the smallest $\chi^2$ is taken as the best-fit redshift of the C1 galaxy, which is 0.1276 with the 1-$\sigma$ error of 0.0004 ($\Delta \chi^2 = 1$). Recalling the redshift of J1634+2049

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16 As the whole profile of [O\textsc{ii}]$\lambda$5007 is blueshifted (see Figure 5), which is not the case of the other unshifted narrow emission lines (e.g., H\textsc{ii}/H\textsc{ii}). We only integrate fluxes within $-500$–500 km s$^{-1}$ corresponding to the H\textsc{ii} narrow lines.
The best redshift is \( z = 0.1293 \), the redshifts of the J1634+2049 and C1 implies that the two galaxies form a collisional system with a line of sight (LOS) velocity difference of \( \approx 451 \text{ km s}^{-1} \). We also calculate the \( V \) band luminosity for C1 by convolving the spectrum with Jonson \( V \) band filter, yielding \( L_V = 7.8 \times 10^{42} \text{ erg s}^{-1} \). Though the best-matched template for C1 shows a stellar age of 5 Gyr, the real stellar age of C1 is hard to determine with the low signal-to-noise ratio. For a stellar population of solar metallicity with an age of 5 Gyr, the real stellar age of C1 is older than 9 Gyr. To be conservative, we argue the stellar age of C1 should be older than 5 Gyr. For a stellar population of solar metallicity with an age of 5 Gyr and 12 Gyr, the mass-to-light ratios \( M/L_V \) are between 3.36 and 6.68 \( M_\odot/L_\odot \) (Bruzual & Charlot 2003). Thus, the stellar mass of C1 \( M_{c1} \) is estimated to within \( 7.3 \times 10^9 \text{ to } 1.45 \times 10^{10} M_\odot \).

Figure 2 shows that J1634+2049 is an ETG (an elliptical or a S0 galaxy) with a ring structure. The off-centered ring is an ellipse in the sky, with its major and minor axes being \( \sim 14\arcmin \) (32.3 kpc) and \( \sim 12\arcmin 7 \) (29.4 kpc), respectively. In Section 2.3 we measured the quantities of the ring as well as the main body (namely the Sérsic of the host galaxy from the SDSS images (see Table 2). The colors \( (g - r, r - i) \) of the Sérsic component are \( (0.93, 0.5) \), and the ring, \( (0.94, 0.48) \). The similarity of the colors between the two components suggests that the ring may have the same stellar population as the Sérsic component, (averagely) aging \( \approx 9 \text{ to } 12 \text{ Gyr} \). Taking the above implication, the ring has \( L_{V, \text{ring}} = 1.9 \times 10^{43} \text{ erg s}^{-1} \), and stellar mass \( M_{\text{ring}} = 2.9 \times 10^{10} M_\odot \).

There are several theories proposed to explain the formation and evolution of the ring galaxies: (1) a head-on collisions between a disk galaxy and an intruder with a mass of at least one tenth of the disk galaxy (e.g., Lynds & Toomre 1976; Theys & Spiegel 1977; Toomre 1978, p. 109; Mapelli et al. 2008; Mapelli & Mayer 2012; Smith et al. 2012); (2) Lindblad resonances which form a smooth ring with a central nucleus and with the absence of companion galaxies (Buta et al. 1999); and (3) an accretion scenario which forms the polar ring galaxies (Bournaud & Combes 2003). The ring around J1634+2049 and observation on the nearby galaxy, C1, suggest a head-on collision scenario for the formation of the ring. Besides, the same color of the stars in the ring and the host galaxy confirms such a formation history. The ring around J1634+2049 is offset to the galactic center (the nucleus), which resembles the collisional RN class of galactic rings proposed by Theys & Spiegel (1976). Numerical simulations showed that the ring structure can be created in the stellar, the cold gas, or both contents of galactic disk by the radial propagation of a density wave which is formed in the collision (e.g., Lynds & Toomre 1976; Theys & Spiegel 1977; Toomre 1978, p. 109; Mapelli et al. 2008; Mapelli & Mayer 2012). Numerical simulations also showed that an off-center collision can produce the offset of the central nucleus and the elliptical rings (e.g., Lynds & Toomre 1976; Mapelli & Mayer 2012; Wu & Jiang 2015). The inclination angle of the collisional parent
galaxies are related to the ellipticity of a galactic ring (Wu & Jiang 2015, Equation (7)). The ellipticity of the ring around J1634+2049 is $\epsilon = 1 - b/a \approx 0.09$, where $a$ and $b$ are the major and minor axes of the projected ring. The ellipticity of the ring is small, so that the off-center collision and inclination effects are small. Besides, the impact parameter (i.e., the minimal distance between the bullet and the disk galaxy) sensitively affects the morphology of the ring (Toomre 1978, p. 109). The ring is more lopsided and the nucleus is more offset when the impact parameter is larger. The ring structure disappears when the impact parameter is too large, and spiral structure forms instead. The ring structure is clearly around the J1634+2049, and it is offset. Thus we may infer that J1634+2049 and C1 had experienced an off-centered collision with a small impact parameter.

### 3.4. Determining the Physical Condition of the Outflows

As described in Section 2.4.3, strong blueshifted absorption troughs show in the optical and NIR spectra, indicating the presence of strong outflows. The decomposition of the emission line profiles (Section 2.4.2) also indicates that the presence of outflows in emission as revealed by the blueshifted line components of almost all the observed emission lines. In this subsection, we analyze and determine the physical properties of both the absorption line and emission line outflows using photoionization synthesis code Cloudy (c13.03; Ferland et al. 1998).

#### 3.4.1. The Absorption Line Outflow

J1634+2049 shows He $\lambda\lambda 3889, 10830$ and Na I D absorption lines in its NUV, optical and NIR spectra. The physical conditions of the absorbers are quite different to generate He $\lambda\lambda 3889, 10830$ and Na I D absorption lines. The metastable $2s$ state in the helium triplet, He $\lambda\lambda 3889, 10830$ is populated via recombination of He $^2S$ ions, which is ionized by photons with energies of $h\nu \geq 24.56$ eV. Therefore, He $\lambda\lambda 3889, 10830$ is a high-ionization line and its column density ($N_{\text{He}^+}$) mainly grows in the very front of hydrogen ionization front and stops growing behind it (e.g., Arav et al. 2001; Ji et al. 2015; Liu et al. 2015). Instead, Na I D absorption line is produced by neutral Sodium, the potential of which is only 5.14 eV, and is easily destroyed by the hard AGN continuum; therefore, this line is rare to detect in quasar spectra. It can only exist where dust is present to shield neutral sodium from the intense UV ($\gtrsim 5.14$ eV) radiation filed of the AGN that would otherwise photoionize it to Na $^+$. In Section 2.4.3 we get the column densities of He $\lambda\lambda 3889, 10830$ and Na I D for the major absorption trough A+B (see Figure 7), $\log N_{\text{He}^+} = 15.01 \pm 0.16$, $\log N_{\text{Na}^+} = 13.31 \pm 0.21$. Low-ionized Ca $^{2+}$ H & K absorption lines usually present with Na I D lines. Ca $^{2+}$ H & K absorption lines arise from Ca $^{+}$ ions, which is ionized by photons with energies of $h\nu \geq 6.11$ eV and destroyed by photons with energies of $h\nu \geq 11.8$ eV. We find no apparent Ca $^{2+}$ H & K absorption lines in the NUV spectrum, which is probably because calcium element is depleted into the dust grains. We estimate the upper limit of column densities Ca $^{2+}$ by shifting Na I D absorption profile to Ca $^{2+}$ H wavelength, and use the profile as a template to fit the Ca $^{2+}$ H region. The upper limit of column density of Ca $^{2+}$ is $\log N_{\text{Ca}^{++}} = 12.9$. All of the above indicate that the dust should be considered in the photoionization models. In addition, we also notice there are no apparent Balmer-line absorption lines in the spectrum of J1634+2049, and we get the upper limit of column densities of hydrogen $n = 2$ is $\log N_{\text{H}}(n=2) = 12.6 \text{ cm}^{-2}$. This suggests the electron densities of the absorbing gas cannot be higher than $10^5 \text{ cm}^{-2}$ (Leighly et al. 2011; Ji et al. 2012).

Here we present a simple model calculated by Cloudy (c13.03; Ferland et al. 1998) to explore the conditions required to generate the measured He $\lambda\lambda 3889, 10830$ and Na I D column densities. We start by considering a gas slab illuminated by a quasar with a density of $N_{\text{H}}$ and a total column density of $N_{\text{H}}$. The SED incident on the outflowing gas has important consequences for the ionization and thermal structures within the outflow. Here we adopt the UV-soft SED, which is regarded more realistic for radio-quiet quasars than the other available SEDs provided by Cloudy (see the detailed discussion in Section 4.2 of Dunn et al. 2010) The UV-soft SED we adopted here is a superposition of a blackbody “big bump” and power laws, and is set to be the default parameters given in the Hazy document of Cloudy where $T = 150,000$ K, $\alpha_{\text{ox}} = -1.4$, $\alpha_{\text{uv}} = -0.5$, $\alpha_{\text{x}} = -1$, and the UV bump peaks at around 1 Ryd.

The above analysis shows the absorption material is a mixture of dust grains and gas, so that the dust-to-gas ratio and the depletion of various elements from the gas phase into dust should be taken into account in the models. The data of J1634+2049 in hand favor the presence of dust of the LMC extinction type, which certainly needs UV spectroscopic observations to confirm. In the following model calculations, we use the Cloudy’s built-in model of ISM grains and assume the total (dust+gas) abundance of the absorption material to be the solar abundance. So the gas-phase abundance changes with the dust-to-gas ratio. As the Cloudy software does not account for the conservations of the mass and abundance of the elements that are both in the gas phase and depleted in the dust grains, trying to keep the total abundances of various element to be the solar. The Cloudy’s built-in grain model of ISM dust that incorporates C, Si, O, Mg and Fe elements. Regarding the Ca and Na elements, we adopt the built-in scheme of dust depletion, with their gas-phase depletion factors being $10^{-4}$ and 0.2, respectively, which are the default recipe of Cloudy. Besides element abundances, the dust-to-gas ratio ($A_{V}/N_{\text{H}}$) is another key parameter in the Cloudy model involving dust. According to the measured narrow-line and broad-line extinction, we simply adopt $A_{V} = 2$ for the dust in the cloud slab. The total column density $N_{\text{H}}$ of the slab can be mutually constrained by comparing the measured Na I D, He $\lambda\lambda 3889, 10830$ and the non-detections (upper limits) of Ca I and Balmer absorption lines with the Cloudy simulations in the parameter spaces.

We start the simulations first with the dust-free baseline models to get the initial total column density $N_{\text{H}}$ (the most probable value; see below), and then we feed Cloudy with the dust-to-gas ratio, $A_{V}/N_{\text{H}}$ calculated by this $N_{\text{H}}$ and get a new $N_{\text{H}}$. Then we use this new $N_{\text{H}}$ into Cloudy and start a new iteration. The simulations will be stopped when the $N_{\text{H}}$ value is convergent. During every iteration, the Cloudy simulations are run over the grids of the following parameter space: ionization parameter $2.5 \leq \log U \leq 1.0$, hydrogen density $3 \leq \log n_{\text{H}}$ (cm$^{-3}$) $\leq 7$, and the stop column density $21 \leq \log N_{\text{H}}$(cm$^{-2}$) $\leq 25$ (see Liu et al. 2015 for the detail of the Cloudy

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17 The full description of the depletion scheme is given in Section 7.9.2 of the Cloudy 13.03 manual Hazy and references therein.
modeling). By comparing in the $N_H-U$ plane the measured ionic column densities with those simulated by Cloudy, we get the allowed parameter intervals for $N_H$ and $U$ as constrained mutually by the various absorption lines. The details are illustrated in Figure 10.

In Figure 10 the upper three panels display dust-free models and the lower three panels display the best models added the effects of dust grains mixed in the gas slab. The red and the blue dashed lines represent the contours of ionic column densities of He I* and Na I, respectively. The red and the blue shaded regions represent the locus of points ($N_H, U$) able to reproduce the observed $N_{He I*}$ and $N_{Na I}$, with 1-σ error, respectively. Purple and the green dotted-dashed lines represent a upper limit on the ionic column densities of He($n = 2$) and Ca II. Correspondingly, the purple and the green dotted lines represent the upper limit minus 3-σ error of $N_{He I*}$ and Ca II, which indicate the confidence interval of the upper limit and also the decreasing direction of the ionic column density. The best models are marked by an open square with the systemic on the solution.

### 3.4.2. Emission-line Outflows

As Figure 5 demonstrates, J1634+2049 also shows emission line outflows as revealed by blueshifted hydrogen Balmer and Paschen lines, He I* λ10830, 5876, [O III]λλ4959, 5007, [O II] λ3727 and Ne III λ3869 lines, particularly by the blueshifted, well-separated cuspy components. This narrow (FWHM $\lesssim 650$ km s$^{-1}$ corrected for the instrumental broadening) cusp exists only in the recombination lines, invisible in any forbidden emission lines. It is blueshifted by 900 km s$^{-1}$ with respect to the system redshift consistently in the hydrogen Balmer and Paschen lines and He I lines; particularly, in the H$\beta$ and He λ10830 this component is well-separated from the peaks of the normal NLR- and BLR-originated components, obviously not an artifact of the line profile decomposition. Interestingly to note, in light of its absence in forbidden lines, this line component should be originated from the dense part of the outflows. A second blueshifted line component, is present
in all high-ionization emission lines such as He I $\lambda$5876, 10830, [O III] $\lambda$4959, 5007 and Ne m 33869, with a similar asymmetric profile and similar best-fit FWHM of $\approx$1600 km s$^{-1}$. This component is much stronger in flux and relatively broader than the cuspy component. Although in He I 10834 it is not so well-separated from the normal NLR and BLR components, this component manifests itself well in the forbidden lines [O III] $\lambda$4959, 5007 and Ne m 33869 with the whole emission line profile being blueshifted, since there are no other line components in these forbidden lines. The low-ionization forbidden line [O I] $\lambda$3727 also shows an obvious bulk-blueshifted component, in addition to the normal NLR-originated component sit just at the system redshift that exists in all low-ionization, forbidden emission lines (e.g., [N II] and [S II]); this bulk-blueshifted component of [O I] $\lambda$3727 has a smaller blueshift than in the aforementioned high-ionization lines. Besides, we can infer that this second blueshifted component (denoted as the broad outflow) should be originated from the less dense part of the outflows, with density lower than $10^4$ cm$^{-3}$.

The dynamics of the outflowing gas of this object should be complex. Here we only consider about the strong, blueshifted broad component of the emission lines, which is much stronger than the cuspy narrow component. To investigate the physical condition for the emission line outflow gas, we use the Cloudy simulations calculated above and then confront these models with the measured line ratios to determine $N_{\text{H}}$, $n_{\text{H}}$, and $U$. As we demonstrate above, both the continuum and emission lines of J1634+2049 are heavily reddened by dust, so it is better for us to use the line ratios of adjacent lines to minimize the effect of reddening. Besides, the difference of ionization potential of the two lines should be large in order for them to probe different zone in a gas cloud. Thus, we use the line ratios [O III] $\lambda$5007/H$\beta$ and He I $\lambda$10830/P$\gamma$ here. The flux of [O III] $\lambda$5007 is measured by subtracting the flux within $\pm$500 km s$^{-1}$ (see Footnote 8) from the whole [O III] $\lambda$5007 profile, i.e., only the flux blueward of $-500$ km s$^{-1}$; the flux of He I $\lambda$10830 is the broad blueshifted component (see Table 3). We do not detect the broad blueshifted component in hydrogen emission lines, which may be weak and concealed in the best-fit blueshifted narrow component; so we take the blueshifted narrow components of H$\beta$ and P$\gamma$ as the upper limits here. Therefore, the lower limits of [O III] $\lambda$5007/H$\beta$ and He I $\lambda$10830/P$\gamma$ are estimated to be 12.75 $\pm$ 1.12 and 15.28 $\pm$ 3.44. On the modeling part, we first extract the simulated [O III] $\lambda$5007, He I $\lambda$10830, H$\beta$, and P$\gamma$ fluxes from the dust-free Cloudy simulations and then compute the line ratios [O III] $\lambda$5007/H$\beta$ and He I $\lambda$10830/P$\gamma$. Figure 11 shows results of the dust-free models. The violet and green lines show the observed line ratios of He I $\lambda$10830/P$\gamma$ and [O III] $\lambda$5007/H$\beta$, which are lower limits of the actual line ratios for the outflow gas (see above). So the region enclosed by the violet red and green lines is the possible parameter space for the outflow gas of this object. Model calculations suggest that gas clouds with log $19 < N_{\text{H}}$(cm$^{-2}$) $< 23$ can generate the observed line ratios of this object (Figure 11). Both He I $\lambda$10830/P$\gamma$ and [O III] $\lambda$5007/H$\beta$ are sensitive to the hydrogen front, so the gas which is too thin to contain a hydrogen front or gas which is too thick cannot generate the observed line ratios. Although the appropriate $N_{\text{H}}$ spread 6 dex, the densities $n_{\text{H}}$ are confined to $4 \leq \log n_{\text{H}}$(cm$^{-3}$) $\leq 6$, which is very close to the condition of absorption line outflow gas. Thus, we infer that the blueshifted emission lines are produced in the same outflowing material as the BALs. Based on this assumption, we extract these line ratios from the convergent dusty models (see the lower three panels of Figure 10). In Figure 12 we only shows models with $N_{\text{H}}$ = $10^{22.5}$ cm$^{-2}$ and $N_{\text{H}}$ = $10^{23}$ cm$^{-2}$, which cover the parameter space of the convergent dusty models. The suitable parameter space to generate the observed line ratios of J1634+2049 is $3.5 \lesssim \log n_{\text{H}}$(cm$^{-3}$) $\lesssim 5.5$ and $-1.8 \lesssim \log U$ $\lesssim -0.7$, which is well consistent with the results of absorption line outflows. Combined with the gas parameters determined from the absorption lines, the acceptable parameters ($\log n_{\text{H}}$, log $N_{\text{H}}$, log $U$) of the outflow are $4 \leq \log n_{\text{H}} \leq 5$, $22.5 \leq \log N_{\text{H}} \leq 22.9$, and $-1.3 \leq \log U \leq -0.7$. Note that the grids with $\log n_{\text{H}} = 4$ cm$^{-3}$ is on the edge of the allowed parameter space.

3.4.3. Kinetic Luminosity and Mass Flux of the Outflow

After analyzing the absorption line and emission line outflows, we find that the hydrogen density of both is quite similar and that the derived values of the ionization parameter $U$ are also consistent. This suggests that the blueshifted emission lines are plausibly originated in the absorption line outflows. If this is the case, we can accurately determine the kinetic properties of the outflows by taking advantages of both the absorption lines and the emission lines. Specifically, the absorption lines, which trace the properties of the outflow in the LOS, is good at determining the total column density ($N_{\text{H}}$) and the velocity (gradient) of the outflow; the emission lines, which trace the global properties of the outflows, can determine better the total mass and global covering factor of the outflows.

As the first step, we determine the distance ($R$) of the outflow (exactly speaking, the part producing the LOS absorption) away from the central source. The ionization parameter $U$ depends on $R$ and hydrogen-ionizing photons emitted by the central source ($Q_{\text{H}}$), as follows:

$$U = \frac{Q_{\text{H}}}{4\pi R^2 n_{\text{H}} c},$$

where $n_{\text{H}}$ is the density of the outflow and $c$ is the speed of light. The $n_{\text{H}}$ has been estimated as $\log n_{\text{H}} \approx 4-5$ cm$^{-3}$. To determine the $Q_{\text{H}}$, we scale the UV-soft SED to the de-reddened flux of J1634+2049 at the WISE W1 band ($\sim$3.4 $\mu$m) (see Figure 13), and then integrate over the energy range $h\nu \geq 13.6$ eV of this scaled SED. This yields $Q_{\text{H}} = 4.2 \times 10^{55}$ photons s$^{-1}$. To check the reliability of this integration, we also integrate this scaled SED over the whole energy range and get $L_{\text{bol}} = 4.3 \times 10^{45}$ erg s$^{-1}$, which is basically consistent with our estimated $L_{\text{bol}}$ independently from the obscuration-corrected continuum luminosity at 5100 Å (see Section 3.1). Using this $Q_{\text{H}}$ value together with the derived $n_{\text{H}}$ and $U$, the $R$ value can be derived, as listed in Table 6. With $R$ being 48–65 pc, the outflow is located exterior to the torus while the extend of the torus is on the scale of $\sim$10 pc (Burtcher et al. 2013). This result is consistent with our qualitative analysis of the normalization for the intrinsic spectrum underlying the absorption trough (see Section 2.4.3).

Assuming that the absorbing material can be described as a thin ($\Delta R/R \ll 1$) partially filled shell, the mass-flux rate ($M$)
and kinetic luminosity \( \dot{E}_k \) can be derived as follows (see the discussion in Borguet et al. 2012):

\[
M = 4\pi R\Omega \mu m_p N_H v, \quad (6)
\]

\[
\dot{E}_k = 2\pi R\Omega \mu m_p N_H v^3, \quad (7)
\]

where \( R \) is the distance of the outflow from the central source, \( \Omega \) is the global covering fraction of the outflow, \( \mu = 1.4 \) is the mean atomic mass per proton, \( m_p \) is the mass of proton, \( N_H \) is the total hydrogen column density of the outflow gas, and \( v \) is the weight-averaged velocity of the absorption trough, which is directly derived from the trough’s profile. The weight-averaged \( v \) for the He\(^+\) 3889 absorption trough is \(-3837\) km s\(^{-1}\). Note that the outflow velocity \( v \) should be calculated with the absorption line velocity, not with the outflowing emission line velocity. The absorption lines are produced from the absorber moving along our LOS, whereas the emission lines originate from gas outflowing along different directions with respect to the observer. Thus, the observed outflow velocity of an emission line is a sum of the projected velocities of the outflowing gas along different directions and should be smaller than the outflow velocity of the absorbing material; this is just as we observed.

We estimated the global covering fraction \( \Omega \) for J1634+2049 by comparing the measured EW([O\textsc{iii}] 5007) with the predicted one by the best Cloudy model (see Section 3.4.2). Although to this end, theoretically it is better to use recombination lines such as H\(\beta\) and He\(\lambda 10830\); there is, however, no good measurements of the H\(\beta\) for the relatively broad, blueshifted component (see discussion in Section 3.4.2) and that component of He\(\lambda 10830\) is heavily affected by the absorption trough. The EW([O\textsc{iii}] 5007) value is affected by the dust extinction both to the continuum and the [O\textsc{iii}] 5007 emission. We make simple and reasonable assumptions as follows (cf. Section 2.4.2): the continuum suffers dust extinction.

Figure 11. Photoionization models of the emission line outflow in J1634+2049 assuming solar abundances and dust free. Different panels are for models with different \( N_H \). The green and violet dashed lines show the contours of line ratios of [O\textsc{iii}] 5007/H\(\beta\) and He\(\lambda 10830\)/P\(\gamma\) in \((U, N_H)\), respectively. The green and violet solid lines represent the locus of the measured upper limit for line ratios of [O\textsc{iii}] 5007/H\(\beta\) and He\(\lambda 10830\)/P\(\gamma\), respectively. The closed region surrounded by the green and violet solid lines are the possible parameter space for the emission line outflow of J1634+2049.
Figure 12. Dust models of the emission line outflow in J1634+2049. The emission line ratios are directly extracted from the best dust models of absorption line outflow. The symbols are the same as Figure 11.

Figure 13. Observed spectrum (the black solid lines) and photometry data (the red dots) of J1634+2049. The green solid curve represents the UV-soft SED scaled to the observed flux at the WISE W1 band. The green dashed curve shows the UV-soft SED reddened with the LMC extinction curve. The gray dotted line is the average QSO spectrum as shown in Figure 1.
extinction to the same degree of the broad lines with $E_B - V = 0.64$, and the [O II] λ5007 emission, within the range of dust-free and the broad-line one. After being corrected for the dust extinction, the actual EW([O II]) should be in the range of 4.3–24.8 Å. Here the outflowing [O II] flux is the same as used in Section 3.4.2 and the continuum flux is measured from the decomposed power-law component at 5007 Å. In Cloudy, the emerging [O II] flux is out put with the covering fraction assumed to be 1. The derived EW([O II]) is 82 Å for the model with $n = 10^{4.5}$ cm$^{-3}$, and 142 Å for $n = 10^{5}$ cm$^{-3}$. Thus, the global cover fraction ($\Omega$) for J1634 +2049 is estimated to be in the range of 5.2%–30.1% for $n = 10^{4.5}$ cm$^{-3}$, or 3.0%–17.4% for $n = 10^{5}$ cm$^{-3}$. Likewise, we estimate the $\Omega$ for the outflow emitting He λ10830, yielding 72.9%–100% (the case of $n = 10^{4.5}$ cm$^{-3}$) or 43.7%–71.7% ($n = 10^{5}$ cm$^{-3}$), which are much larger than those for [O II] λ5007. The large difference between the $\Omega$ values estimated based on He λ10830 and [O II] may be mainly due to the measurement uncertainty of the outflowing component of He λ10830, and/or may reflect the inhomogeneity of the outflowing gas. [O II] is a forbidden line that traces the region of low density only and He λ10830, a recombination line, can be generated in much broader spatial regions. Here we conservatively adopt the $\Omega$ value estimated from [O II]. Thus, the kinetic luminosity and mass loss rate are calculated as summarized in Table 6.

4. DISCUSSION AND SUMMARY

We performed a comprehensive multiwavelength study of the properties of J1634 +2049, a local quasar with high IR luminosity of $L_{IR} = 10^{11.91}$ $L_\odot$, and signatures of outflows both in BALs and blueshifted emission lines. The high IR luminosity indicates recent violent star formation activities with SFR = 140 ± 43 $M_\odot$ yr$^{-1}$, which is also indicated by the P AH emission lines with a derived SFR = 141 ± 71 $M_\odot$ yr$^{-1}$. Interestingly, in the UV and optical bands there are few signs of star formation activities, and the decomposition of the SDSS images demonstrate the host galaxy is of early type in terms of its structural property (e.g., Sérsic $n = 4$) and the colors (e.g., $g - r = 0.93$. This quasar has a circumgalactic ring on scales of ~30 kpc. The ring has almost the same colors as the host galaxy (namely the Sérsic $n = 4$ component). Yet it is unclear at this point whether the galactic ring has a similar intense dust-enshrouded star formation. There are two small galaxies to the west of it within 20″ on the sky, the three of them being in an almost linear configuration. We spectroscopically observed the galaxies and obtained the redshift for the relatively bigger one (C1; the closer to J1634 +2049), $z = 0.129^{+0.007}_{-0.005}$.

In the optical and near-infrared spectra, there are several strong BALs and blueshifted emission lines in addition to the normal, BLR- and NLR-originated emission lines, suggestive of the AGN outflows. These lines can be used to derive (mutually constrain) the physical properties of the outflowing gas by confronting the observed with the modeling results of the photoionization software Cloudy. The appearance of the BAL of neutral Sodium, which is rare in an AGN environment, suggests that the outflowing gas is thick and dusty. The physical parameters we determined with Cloudy for absorption line and emission line outflows are very close, with $10^{4.5} \lesssim n_H \lesssim 10^5$ cm$^{-3}$, $10^{-1.3} \lesssim U \lesssim 10^{-1.0}$, and $N_{HI} \approx 10^{22.5}$ cm$^{-2}$. This similarity suggests that those absorption and emission lines should be generated in the common outflowing gas. Using the absorption lines to derive the total column density of the outflow and the emission lines to obtain the global covering factor, we estimate the distance of the absorbing material to the central source, $R \approx 48–65$ pc, exterior to the torus. These derived parameters allow us to calculate the kinetic luminosity and mass loss rate of the outflow. The results listed in Table 6 indicate that the outflow observed in J1634 +2049 processes a large kinetic luminosity, which is high enough to play a major role in the AGN feedback. Previous studies suggest that AGN feedback typically requires a mechanical energy input of roughly 0.5%–5% of Eddington luminosity of the quasar to heat the cold gas and quench the star formation activities in the host galaxy (e.g., Scannapieco & Oh 2004; Hopkins & Elvis 2010).

Taking the multiwavelength results altogether, we can infer the whole story as follows. J1634 +2049 could have been a disk galaxy with abundant gas, and was collided through by one or two small galaxies with masses of ~1 × 10$^{10}$ $M_\odot$. The violent head-on collision destroyed the disk and formed a circumgalactic ring. The collision also triggered violent star formation in the host galaxy, with the star formation regions being heavily reddened by dust now and betraying itself in the IR bands only. We may infer that the violent star formation is in the circumnuclear region, as the optical images and spectrum demonstrate that the large scale of the host galaxy is an elliptical galaxy dominant by an old stellar population. Indeed, in the generally believed scenario of galaxy merger/collision and AGN feedback, as described in the Introduction section, the large-scale cold gas of the interacting galaxies loses angular momentum and is driven into the nuclear region due to the gravitational potential well; the infall of the cold gas triggers nuclear star formation and also feeds the central SMBH. This inference is further supported by the fact that the narrow emission lines (dominated by the AGN according to the BPT diagrams) as well as the broad emission lines are seriously

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**Table 6**

Physical Properties of the Outflow

| $log n_H$ | $log N_H$ | $log U$ | $v$ | $R$ | $Ω$ | $log L_{bol}$ | $M$ | $log E_k$ | $E_k/L_{bol}$ |
|-----------|-----------|---------|-----|-----|-----|---------------|-----|-----------|----------------|
| $cm^{-3}$ | $cm^{-2}$ | (km s$^{-1}$) | (pc) | (%) | (erg s$^{-1}$) | (M$_\odot$ yr$^{-1}$) | (erg s$^{-1}$) | (%) |
| 4.0$^a$   | 22.85     | −0.55   | −3837 | 64.1 | 18.1–100 | 45.6 | 455–2513 | 45.3–46.0 | 49.2–272 |
| 4.5       | 22.6      | −1.0    | −3837 | 60.6 | 5.2–30.1 | 45.6 | 70–402   | 44.5–45.3 | 7.5–43.5 |
| 5.0       | 22.5      | −1.3    | −3837 | 48.1 | 3.0–17.4 | 45.6 | 25–147   | 44.1–44.8 | 2.8–15.9 |

**Notes.**

$^a$ $L_{bol}$ here is calculated from UV-soft SED used in CLOUDY models, which corresponds to the estimated hydrogen-ionizing photons ($Q_H$) used to estimate the distance of outflows. The $L_{bol}$ in this table is basically consistent with the $L_{bol}$ calculated from $λ_{5100}(5100)$ Å.

$^b$ Grid with $n_H = 10^4$ cm$^{-3}$ is on the edge of the acceptable parameter space.

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obscured by dust, hinting at a common origin of the dust to the optical emission of both the AGN and star formation. The existence of neutral sodium BAL also indicates that the nuclear region is dusty. With huge kinetic luminosity, the outflows launched by the AGN in turn will blow away the gas and dust around the nucleus and even inhibit the star formation in the host galaxy soon or later, as the co-evolution scenario suggests. At this time point, J1634+2049 is just like the prototypical QSO/LIRG composite object, Mrk 231, in the transitional phase emerging out of the dust. According to the differences in both the degree of dust extinction and the mid- and far-infrared SED shape between J1634+2049 and Mrk 231, J1634+2049 should be at the phase immediately after Mrk 231 in the co-evolutionary sequence. So it is a rare object for us to detailly and quantitatively study (or test) the co-evolution scenario of galaxy and SMBH. In particular, taking the advantage of its nearness, we will carry out spatially resolved observations of the cold gas distribution and kinematics on the circumgalactic scale (by, e.g., JVLA), or even on the scale within the host galaxy (by, e.g., ALMA), to directly investigate the cold gas— the vital starring actor playing in the (co-)evolutionary scenario.

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