[C II] LINE EMISSION IN MASSIVE STAR-FORMING GALAXIES AT $z = 4.7$

J. Wagg$^1$, T. Wiklind$^2$, C. L. Carilli$^{3,4,10}$, D. Espada$^5$, A. Peck$^2$, D. Riechers$^6$, F. Walter$^{7,8}$, A. Wootten$^8$, M. Arauena$^1$, D. Barkats$^2$, J. R. Cortes$^4$, R. Hills$^2$, J. Hodge$^{2,7}$, C. M. Impellizzeri$^{2,8}$, D. Iono$^5$, A. Leroy$^8$, S. Martín$^1$, M. G. Rawlings$^2$, R. Maiolino$^4$, R. G. McMahon$^9$, K. S. Scott$^8$, E. Villard$^5$, and C. Vlahakis$^2$

$^1$ European Southern Observatory, Casilla 19001, Santiago, Chile; jwagg@eso.org
$^2$ Joint ALMA Observatory, Alonso de Cordova 3107, Vitacura, Santiago, Chile
$^3$ National Radio Astronomy Observatory, Socorro, NM, USA
$^4$ Cavendish Laboratory, University of Cambridge, Cambridge, UK
$^5$ National Astronomical Observatory of Japan, Tokyo, Japan
$^6$ Department of Astronomy, California Institute of Technology, Pasadena, CA, USA
$^7$ Max-Planck Institute for Astronomy, Heidelberg, Germany
$^8$ National Radio Astronomy Observatory, Charlottesville, VA, USA
$^9$ Institute of Astronomy, University of Cambridge, Cambridge, UK

Received 2012 April 24; accepted 2012 May 16; published 2012 June 4

ABSTRACT

We present Atacama Large Millimeter/submillimeter Array (ALMA) observations of the [C II] 157.7 μm fine structure line and thermal dust continuum emission from a pair of gas-rich galaxies at $z = 4.7$, BR1202-0725. This system consists of a luminous quasar host galaxy and a bright submillimeter galaxy (SMG), while a fainter star-forming galaxy is also spatially coincident within a 4″ (25 kpc) region. All three galaxies are detected in the submillimeter continuum, indicating FIR luminosities in excess of $10^{13} L_\odot$ for the two most luminous objects. The SMG and the quasar host galaxy are both detected in [C II] line emission with luminosities $L_{\text{[CII]}} = (10.0 \pm 1.5) \times 10^9 L_\odot$ and $L_{\text{[CII]}} = (6.5 \pm 1.0) \times 10^9 L_\odot$, respectively. We estimate a luminosity ratio $L_{\text{[CII]}} / L_{\text{FIR}} = (3.8 \pm 1.2) \times 10^{-4}$ for the starburst SMG to the north and $L_{\text{[CII]}} / L_{\text{FIR}} = (2.5 \pm 0.4) \times 10^{-4}$ for the quasar host galaxy, in agreement with previous high-redshift studies that suggest lower [C II]-to-FIR luminosity ratios in quasars than in starburst galaxies. The third fainter object with a flux density $S_{\text{SMG/HZ}} = 1.9 \pm 0.3$ mJy is coincident with a Lyα emitter and is detected in HST ACS F775W and F814W images but has no clear counterpart in the $H$ band. Even if this third companion does not lie at a redshift similar to BR1202-0725, the quasar and the SMG represent an overdensity of massive, infrared luminous star-forming galaxies within 1.3 Gyr of the big bang.

Key words: cosmology: observations – early universe – galaxies: evolution – galaxies: formation – galaxies: high-redshift

Online-only material: color figure

1. INTRODUCTION

The formation of massive galaxies at high redshift is often traced by intense emission at observed submillimeter (submm) and millimeter (mm) wavelengths, an indication of dust heated by star formation or an active galactic nucleus (AGN). The observed populations of sub-mm-selected starburst galaxies (SMGs; e.g., Blain et al. 2002) and far-infrared (FIR) luminous quasar host galaxies (e.g., Omont et al. 1996; Wang et al. 2008) likely represent the formation of the most massive galaxies at high redshift given their extreme FIR luminosities.\footnote{10} $L_{\text{FIR}} \sim 10^{13} L_\odot$. About one-third of optically luminous quasar host galaxies at $z \gtrsim 5$ display such high FIR luminosities (e.g., Wang et al. 2008), and the implied star formation rates (SFRs) $> 10^9 M_\odot$ yr$^{-1}$ may be the main source of power. If these rates can be maintained for greater than $10^7$ years, such objects could evolve into massive early-type galaxies with old stellar populations (e.g., Daddi et al. 2009).

BR1202-0725 was discovered at $z = 4.7$ with radio continuum emission resolved into two distinct components (Omont et al. 1996; Yun et al. 2000; Carilli et al. 2002), where the southern emission is associated with an optically luminous quasar, and the northern a heavily obscured starburst SMG. These are separated by 3′.8 (25 kpc) in the northwest direction, and each is detected in interferometric imaging of the high-J CO line emission (Omont et al. 1996; Guilloteau et al. 1999). The single-dish CO $J = 1–0$ line luminosity indicates a combined mass of $\sim 8 \times 10^{10} M_\odot$ in cold molecular gas (Riechers et al. 2006). The difference in the high-J CO line redshifts and line profiles demonstrates that BR1202-0725 is not gravitationally lensed, but rather the ongoing merger between two gas-rich galaxies. Submm- and mm-wavelength imaging with the Submillimeter Array (SMA) and the Plateau de Bure Interferometer (PdBI) resolves the thermal dust continuum emission into the northern and southern components (Omont et al. 1996; Iono et al. 2006), revealing extreme FIR luminosities and an estimated SFR $> 1000 M_\odot$ yr$^{-1}$ for each component. The likely interaction between two massive galaxies with such intense multi-wavelength emission makes this an ideal system for studying the formation of galaxies in the young universe.

The brightest emission line in the interstellar medium of galaxies is typically the 157.7 μm [C II] line, arising primarily in photodissociation regions (PDRs) and the cold neutral medium. Its intensity and ubiquitous nature make this line an ideal means to study the formation and kinematics of high-redshift galaxies. The [C II] line luminosity typically contains $\sim 0.1\%–1\%$ of the total FIR luminosity in quiescent galaxies, and less than 0.1% in ultraluminous infrared galaxies (ULIRGs; Crawford et al. 1985; Stacey et al. 1991; Wright et al. 1991; Malhotra et al. 1997). For galaxies at redshifts $z \gtrsim 1$, this line can be redshifted into the submm-wavelength atmospheric window.
which the target was above 60°.

The total on-source observing time was 25 minutes during the same program (on 2012 January 14) contain full complex gain solutions, resulting in a higher dynamic range and so the data are not included in this Letter. However, these observations were used to verify the flux scale of the data presented here.

Data analysis was performed using the CASA software package. The flux density scale was derived using observations of Titan. We derive a flux density of 17 Jy for the complex gain calibrator, 3C279, located 13° from BR1202-0725. Some uncertainty in the flux density of 3C279 is present due to the lack of system temperature measurements with these data. We estimate this uncertainty to be 15% and include this in our measured flux densities and integrated intensities. A 7 minute calibration cycle was used and the weather conditions were dry (precipitable water vapor ~0.7 mm). The gains measured for 3C279 were stable to within a few degrees and a few percent during the observation. The spectral bandpass shape was also measured using 3C279.

The sensitivity is high enough to allow self-calibration using the line-free emission from BR1202-0725. Amplitude and phase self-calibration were done using a 3 minute averaging time for complex gain solutions, resulting in a higher dynamic range in both the line and continuum images. A continuum map of BR1202-0725 was made by imaging the last three spectral windows, which results in a 340 GHz continuum sensitivity of 0.2 mJy beam⁻¹. A UV plane model for the continuum emission is derived from these spectral windows and subtracted from the first window. After continuum subtraction and resampling to ~28 km s⁻¹ resolution, the channel rms is 2.0 mJy beam⁻¹.

3. RESULTS AND ANALYSIS

The synthesized beam size of the 340 GHz continuum observations is 1.30 × 0.96, corresponding to physical scales of 8.6 × 5.7 kpc at z = 4.7. The total 340 GHz flux density of the system is 43 ± 6 mJy (including calibration uncertainties), in excellent agreement with the previous ~14″ resolution 850 μm SCUBA observations (S_{850μm} = 42 ± 2 mJy; McMahon et al. 1999). Figure 1 shows a map of the 340 GHz continuum emission in BR1202-0725, while Figure 2 shows this thermal dust continuum emission overlaid on the HST ACS F775W, WFPC2 F814W, and NIC2 F160W images. Both the quasar host galaxy to the south and the optically obscured starburst SMG to the north are clearly detected at 340 GHz (Table 1). Gaussian fitting of the continuum emission in the image plane suggests that neither of these sources is resolved. A third faint source is detected in 340 GHz continuum emission to the southwest of the quasar with a flux density S_{340GHz} = 1.9 ± 0.3 mJy. We hereafter refer to this source as BR1202-0725 SW, and its faintness places it below the detection threshold of previous interferometer observations at these frequencies (Iono et al. 2006).

For the quasar and the SMG we adopt the FIR luminosity estimates of Iono et al. (2006), who fit the FIR-to-radio observations using the templates of Yun & Carilli (2002) and also measure similar continuum flux densities for the two components at these frequencies. For BR1202-0725 North, they estimate L_{FIR} = 1.2 × 10^{13} L_⊙, and L_{FIR} = 2.6 × 10^{13} L_⊙ for BR1202-0725 South. Using our measurements of the 340 GHz continuum flux densities and assuming β = 1.5 with dust temperatures of 35 K for the SMG (e.g., Kovács et al. 2006) and 60 K for the quasar (e.g., Beelen et al. 2006; Weiß et al. 2007), we estimate FIR luminosities similar to those estimated...
The results are given in Table 1. The line width measured in the 340 GHz continuum emission measured by ALMA, where the lowest contour level is 0.9 mJy beam$^{-1}$. The synthesized beam size is 1"$^\prime\prime$ × 0"$^\prime\prime$.86. The QSO BR1202-0725 is clearly visible in all three HST images. The two Ly$\alpha$ galaxies detected by Hu et al. (1996) can be seen as faint emission in the F775W and F814W images. Only the NW Ly$\alpha$ galaxy is seen in the F160W image. Both the NW and SW Ly$\alpha$ galaxies are offset from the peaks in the submm continuum emission and are apparently extended toward the QSO.

(A color version of this figure is available in the online journal.)

Table 1

| Source              | $S_{340\text{GHz}}$ (mJy) | $z_{[\text{CII}]}$ | $\Delta V_{\text{FWHM}}$ (km s$^{-1}$) | $S_dV$ (Jy km s$^{-1}$) | $L_{[\text{CII}]}$/$L_{\text{FIR}}$ | $L_{[\text{CII}]}$/$L_{\text{FIR}}$ |
|---------------------|--------------------------|-------------------|---------------------------------------|-------------------------|----------------------------------|----------------------------------|
| BR1202-0725 North   | 18.8 ± 2.8               | 4.6906            | 722 ± 12                              | 14.7 ± 2.2              | 10.0 ± 1.5                      | 8.3 ± 1.2                       |
| BR1202-0725 South   | 18.0 ± 2.7               | 4.6943            | 328 ± 6                               | 9.6 ± 1.5               | 6.5 ± 1.0                       | 2.5 ± 0.4                       |
| BR1202-0725 SW      | 1.9 ± 0.3                | ...               | ...                                   | ...                     | ...                             | ...                             |

Note. * $z_{[\text{CII}]}$ and line width ($\Delta V_{\text{FWHM}}$) are determined from the best-fit parameters of a Gaussian fit to the spectra.

by Iono et al. (2006) and simultaneously reproduce the combined 350 $\mu$m flux density of the pair (Benford et al. 1999). BR1202-0725 SW is not detected in the 1.4 GHz observations of Carilli et al. (2002), implying a 3$\sigma$ upper limit of 48 $\mu$Jy at 1.4 GHz. The dust temperature is uncertain, and so to estimate the FIR luminosity of BR1202-0725 SW we scale down the luminosities of the two bright sources by the ratio of the line flux densities, implying $L_{\text{FIR}} \sim 10^{12} L_\odot$. Similar luminosities are estimated for BzK-selected star-forming galaxies and dust obscured galaxies at $z \sim 2$ (e.g., Pope et al. 2008; Daddi et al. 2008) and we might expect comparable 350 GHz flux densities to arise from a range of different high-redshift galaxy populations.

From the spectral line data cubes we extract spectra at the positions of peak intensity in the submm image. Figure 3 shows the spectra of the three continuum sources where redshifted [CII] line emission is clearly detected in both the quasar host galaxy and the SMG. The peak in the [CII] line emission from both objects is spatially coincident with the FIR continuum and previously observed CO $J = 2 \rightarrow 1$ line emission (Carilli et al. 2002). There is also a hint of line emission at the red end of the spectrum observed toward BR1202-0725 SW, although we cannot exclude the possibility that this emission is spurious. We fit Gaussian line profiles to the two detected [CII] lines and the results are given in Table 1. The line width measured in the quasar host galaxy ($\Delta V_{\text{FWHM}} \sim 330$ km s$^{-1}$) is a factor of $\sim 2$ smaller than that observed in the SMG ($\Delta V_{\text{FWHM}} \sim 720$ km s$^{-1}$), which is in broad agreement with the difference observed in CO emission lines detected in two of these populations at high redshift (e.g., Solomon & Vanden Bout 2005; Carilli & Wang 2006). The line widths are also in excellent agreement with previous observations of CO line emission (Omont et al. 1996; Carilli et al. 2002). The [CII] line width and redshift of BR1202-0725 North are in agreement with the previous SMA observations of Iono et al. (2006).

Following the definition of line luminosity (see Solomon & Vanden Bout 2005), we calculate $L_{[\text{CII}]} = (10.0 \pm 1.5) \times 10^9 L_\odot$ and $L_{[\text{CII}]} = (6.5 \pm 1.0) \times 10^9 L_\odot$ for BR1202-0725 North (SMG) and South, respectively. Combined with our estimates of the FIR luminosity, we calculate [CII]-to-FIR luminosity ratios $L_{[\text{CII}]}/L_{\text{FIR}} = (8.3 \pm 1.2) \times 10^{-4}$ for the SMG and $L_{[\text{CII}]}/L_{\text{FIR}} = (2.5 \pm 0.4) \times 10^{-4}$ for the quasar host galaxy. These ratios are in agreement with that found in other high-redshift galaxies observed in [CII] line emission (e.g., Stacey et al. 2010; Wagg et al. 2010; Cox et al. 2011).

![Figure 2. Archival HST ACS F775W (prop ID 10417), WFPC2 F814W (prop ID 5975), and NICMOS F160W (prop ID 7266) images of the BR1202-0725 system at $z = 4.69$. Contour levels show the intensity of the 340 GHz continuum emission measured by ALMA, where the lowest contour level is 0.9 mJy beam$^{-1}$. The synthesized beam size is 1"$^\prime\prime$ × 0"$^\prime\prime$.86. The QSO BR1202-0725 is clearly visible in all three HST images. The two Ly$\alpha$ galaxies detected by Hu et al. (1996) can be seen as faint emission in the F775W and F814W images. Only the NW Ly$\alpha$ galaxy is seen in the F160W image. Both the NW and SW Ly$\alpha$ galaxies are offset from the peaks in the submm continuum emission and are apparently extended toward the QSO.](image)

#### 4. DISCUSSION

The extreme FIR luminosities observed in BR1202-0725 North and South would imply obscured SFRs in excess of 1000 $M_\odot$ yr$^{-1}$ (e.g., Kennicutt 1998). The SFR for the fainter SW component is $\sim 10$ times lower due its lower FIR luminosity. As such, this system likely represents the buildup of an extremely massive galaxy, possibly about to undergo a major merger. Chandra detects X-ray emission in BR1202-0725 South and tentatively in BR1202-0725 North (Iono et al. 2006), suggesting that the rapid buildup of the stellar mass in these objects is coeval with the growth of a supermassive black hole (e.g., Hopkins et al. 2006).

Mounting evidence for the presence of large-scale outflows arising from AGNs has emerged from the detections of large-scale emission from ionized gas in high-redshift quasars (e.g., Cano-Díaz et al. 2012) and the presence of P-Cygni profiles and extended wings in the molecular lines observed in nearby ULIRGs (e.g., Sakamoto et al. 2009; Feruglio et al. 2010; Aalto et al. 2012; Sturm et al. 2011). The [CII] line profile of the quasar BR1202-0725 South is described very well by a Gaussian profile; however, a low-intensity, broad line-width extension is seen toward positive velocities. If all of the ionized gas in this galaxy were to arise from a region of gas rotating in virial...
The redshift of this object is unknown. The rms is 1.3 mJy per channel in the spectrum of BR1202-0725 SW.

The luminosity ratio between emission from the [C\textsc{ii}] line and the FIR continuum provides constraints on the physical conditions of the interstellar gas in galaxies. The [C\textsc{ii}] line typically exhibits a luminosity \(~0.1\%–1\%\) that of the total FIR luminosity in quiescent galaxies, and less than 0.1\% in ULIRGs (Crawford et al. 1985; Stacey et al. 1991; Wright et al. 1991; Malhotra et al. 1997). This ratio may also be lower in luminous quasars than in starburst galaxies (e.g., Stacey et al. 2010), which is supported by the observations presented here.

One explanation for the decrease in the [C\textsc{ii}]-to-FIR ratio in ULIRGs is an increase in the rate of UV photons absorbed by dust grains, which decreases the number of photons available to ionize the gas (e.g., Abel et al. 2009; Papadopoulos et al. 2010). Graciá-Carpio et al. (2011) find that the scatter in the [C\textsc{ii}]-to-FIR ratio can be decreased if one plots this ratio as a function of the FIR luminosity weighted by the molecular gas mass. In the case of BR1202-0725 the CO $J = 2–1$ line luminosity (or the cold molecular gas mass) is the same in both the northern and southern sources (Carilli et al. 2002), and so weighting the FIR luminosity by the molecular gas mass would have an equivalent impact on each. Similarly, high signal-to-noise [C\textsc{ii}] line spectra and FIR continuum observations of significant samples of high-redshift quasars and starburst galaxies will reveal the typical structure and kinematics of the ionized gas in these objects, and finally resolve the origin of the observed line to continuum ratios.

The presence of two SMGs in the BR1202-0725 system, potentially situated within 25 kpc of the QSO, would likely represent an extreme merger event at high redshift. The two brightest submm components, one of which is the QSO host galaxy, contain large amounts of molecular gas (\(>10^{10} M_\odot\)), and each component is likely to have a total mass (gas + stars) in excess of \(10^{11} M_\odot\), similar to comparably luminous SMGs at these redshifts (e.g., Riechers et al. 2010). Giallongo et al. (1998) searched for the presence of an overdensity of galaxies in the vicinity of BR1202-0725, finding no evidence for an increase in the number density of optically selected galaxies at these redshifts. However, our deep continuum observations with ALMA suggest that the star formation in galaxies associated with this field is heavily obscured. There are two Ly\(\alpha\)-detected galaxies in the vicinity of BR1202-0725 (Hu et al. 1996, 1997). The brightest Ly\(\alpha\) emitter is associated with the NW SMG, while the fainter Ly\(\alpha\) emitter is associated with the SW component. Both of these galaxies are detected in the HST F775W and F814W images, but only the brighter NW galaxy is seen in the NICMOS F160W image (Figure 2).

The FIR luminosities of the three submm components imply a combined SFR of several \(10^3 M_\odot\) yr\(^{-1}\). Several observational facts suggest that star formation is indeed powering the high luminosities. The rest-frame UV light associated with the NW component has been shown to be the result of star formation rather than reflected QSO light (Fontana et al. 1998). The combination of high inferred SFRs and significant [C\textsc{ii}] line emission is consistent with a scenario whereby this line emission originates in PDR regions associated with star formation. Taken together, it implies that star formation, possibly triggered by gravitational interaction, rather than processes associated with the QSO are powering the UV as well as the FIR emission.

Daddi et al. (2009) estimate that the number density of submm luminous starburst galaxies in the redshift range \(3.5 < z < 6.0\)
is $10^{-5}$ Mpc$^{-3}$. Two such galaxies with likely masses $\gtrsim 10^{11}$ $M_\odot$ found within a projected radius of 25 kpc therefore represent a significant overdensity at $z = 4.7$, possibly associated with a “proto-cluster” region when the universe was only 1.3 Gyr old. The most probable descendant of this interaction will be a massive galaxy (several $10^{11}$ $M_\odot$), with very little gas and dust remaining. Massive and evolved galaxies have been found at high redshifts in several deep surveys (e.g., Cimatti et al. 2004) and could represent the result of similarly violent merging events as we are currently witnessing in the BR1202-0725 system.

5. SUMMARY

We present the first ALMA observations of redshifted [C ii] line and thermal dust continuum emission in high-redshift quasar host galaxies, observing BR1202-0725 at $z = 4.7$ as part of commissioning and science verification. Our continuum image reveals a faint submm companion to the previously known quasar host galaxy and FIR luminous SMG; however, the redshift of this object is unknown. We also measure high signal-to-noise detections of the ionized carbon line in both the northern starburst galaxy and the quasar host galaxy to the south. The line widths agree with previous observations of high-J CO line emission, with a narrower line profile observed in the quasar. The decrease in the [C ii]-to-FIR ratio measured in nearby starburst galaxy and high-redshift ULIRGs is also observed in BR1202-0725, with a higher ratio in the starburst galaxy, in agreement with a higher ratio in the starburst galaxy, in agreement with previous observations of high-J CO line spectrum, and higher angular resolution observations of this object with ALMA will allow us to constrain the spatial extent of this gas and measure the total energy and kinematics.

We thank the referee for helpful suggestions based on the original manuscript. This work was co-funded under the Marie Curie Actions of the European Commission (FP7-COFUND). We thank all those involved in the ALMA project for making these observations possible. This Letter makes use of the following ALMA data: ADS/JAO.ALMA#2011.0.00006.SV. ALMA is a partnership of ESO (representing its member states), NSF (USA), and NINS (Japan), together with NRC (Canada) and NSC and ASIAA (Taiwan), in cooperation with the Republic of Chile. The Joint ALMA Observatory is operated by ESO, AUI/NRAO, and NAOJ. D.R. acknowledges support from NASA through a Spitzer Space Telescope grant. The National Radio Astronomy Observatory is a facility of the National Science Foundation operated under cooperative agreement by Associated Universities, Inc.