Single Transit Candidates from K2: Detection and Period Estimation

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ABSTRACT
Photometric surveys such as Kepler have the precision to identify exoplanet and eclipsing binary candidates from only a single transit. K2, with its 75d campaign duration, is ideally suited to detect significant numbers of single-eclipsing objects. Here we develop a Bayesian transit-fitting tool (“Namaste: An MCMC Analysis of Single Transit Exoplanets”) to extract orbital information from single transit events. We achieve favourable results testing this technique on known Kepler planets, and apply the technique to 7 candidates identified from a targeted search of K2 campaigns 1, 2 and 3. We find EPIC203311200 to host an excellent exoplanet candidate with a period, assuming zero eccentricity, of 540+410−230 days and a radius of 0.51±0.05 RJup. We also find six further transit candidates for which more follow-up is required to determine a planetary origin. Such a technique could be used in the future with TESS, PLATO and ground-based photometric surveys such as NGTS, potentially allowing the detection of planets in reach of confirmation by Gaia.

Key words: keywords – keywords

1 INTRODUCTION
Wide-field survey telescopes have for more than a decade searched for the repeated transits of exoplanets, detecting thousands of candidates and confirmed planets. With photometric precision orders of magnitudes better than on the ground, the Kepler mission (Borucki et al. 2010) has contributed most to this growing field, from sub-Earth-radius worlds (Barclay et al. 2013) to long-period gas giants (Wang et al. 2013).

In 2014 Kepler was repurposed as K2 (Howell et al. 2014) which, due to engineering constraints, observes multiple fields in the ecliptic on 75 day campaigns. Although reduced pointing stability limits the photometric precision of K2, many stars have been observed with precision on the order of 100ppm per half hour cadence. More than 40 planet candidates have so far been detected from the first 3 campaigns of K2 (e.g. Foreman-Mackey et al. 2015), with orbital periods up to 50 days. More than a dozen of these systems have subsequently been validated (e.g. Montet et al. 2015; Armstrong et al. 2015a).

The occurrence rates of transiting planets with periods on the order of months (themselves derived by Kepler, e.g. Fressin et al. (2013)) suggest that a handful of longer-period planets should be detected per K2 campaign. The reduced mission duration of 75 days (compared to 1400 in the primary mission) means that such planets are likely to only transit once. Planets with transit depths over ~1mmag offer the potential of discovery in just a single event. Such signals may then constitute strong planetary candidates with undefined orbital parameters, similar to those detected by microlensing surveys (Bennett & Rhie 1996).

One such planet has previously been detected and confirmed from K2. K2-1 b was initially spotted in a single transit during the best 6.5 days of engineering data (Vanderburg et al. 2015). Subsequent follow-up with both spectroscopy (HARPS) and photometry (MOST) determined that the planet was a 2.5R⊕ Earth on a 9.12d orbit. However, in subsequent campaigns, planets capable of being detected in a single transit are likely to have orbital periods of 40 days or more, making follow-up more challenging.

We show here that for favourable transits, provided the star can be characterised, the orbital period can be estimated from the information contained within a single tran-
From its initial sampler of $\sim$150,000 stars, *Kepler* detected 101 planets and 828 planet candidates with periods longer than 50 days. 53% of these exhibit deep enough transits to allow their detection from a single transit ($>5\sigma$). We would therefore expect, with $\sim$40,000 stars now observed by K2, to detect substantial numbers of such planets.

A simple analysis, using *Kepler* occurrence rates from 0 to 85 days (Fressin et al. 2013) and assuming a flat distribution in $\ln(P)$ beyond 85d, suggests that 15% of FGK stars should have a large Neptune or Jupiter on $<3000$ day orbits. This is similar to the giant planet occurrence rates found by (Mayor et al. 2011), with 14 $\pm$ 2% of FGK stars with planets larger than $>50M_\oplus$. Accounting for transit (scaled with $R_p^2$) and timing probability (scaled with $T_p^2$) gives detectable multi- and single-transiting planets around 0.15% & 0.03% of stars from Fressin et al. (2013) and 0.09% & 0.02% from Mayor et al. (2011). Hence one in every few thousand FGK stars observed by K2 should have detectable single transits.

The decreasing probabilities with orbital distance also suggest that high-precision, shorter-duration surveys such as TESS (Ricker et al. 2015) or NGTS (Wheatley et al. 2013) could detect multi-transits in substantial numbers. The 28d survey duration for the majority of the TESS field, for example, would likely find multi-transiting giant planets around $>0.04%$ FGK stars, potentially improving giant planet yield by as much as 50%.

Occurrence rate estimations of this regime are, at present, extremely poorly constrained. Hence the detection and subsequent follow-up of multi-transiting planets in this regime will allow improved occurrence rates for this relatively unexplored parameter space.

### 2 METHODS

#### 2.1 Detrending

Three campaigns of K2 target pixel files, were obtained from the Mikulski Archive for Space Telescopes (MAST). Limiting our analysis to objects classified as ‘STARS’ we performed aperture photometry on the target object. Aperture sizes were varied according to brightness, with radii of 3, 4, 5 and 12 pixels used for *Kepler* magnitudes bins with boundaries 16, 13 & 10. The extracted flux was then background-subtracted, with background RMS added in quadrature to the flux errors.

To remove the majority of systematics in K2 lightcurves, which are dominated by pointing drift, we independently developed a detrending method similar to (Vanderburg & Johnson 2014) to remove all noise correlated with shifting target position regardless of its source. To do this we compute centroid position for each timestamp, cutting the largest pixel shifts (eg, during thruster firings). We then create a 2D surface of raw flux against x and y position. By binning this to an evenly spaced $10 \times 10$ grid (discarding bins with fewer than 3 points) and interpolating the median fluxes with SciPy’s griddata function (Jones et al. 01 ) we created a smooth surface map of the variation due to centroid shifts. This could then be divided out of the extracted lightcurve, with the result of significantly decreasing the RMS error of the lightcurve. Variations not related to pointing drift (eg, long-timescale flux drifts) are not removed by this technique and we note that, in many cases, some systematic noise remains due to other instrumental effects. This method is described in detail in Armstrong et al. (2015b).\(^1\)

#### 2.2 Transit Search

Long-duration variability was removed from the detrended lightcurves by fitting 3rd order polynomials to 2d windows either side of an untouched 4hr central window. The fit for each was iterated 20 times, with points $>5\sigma$ from the best fit excluded each time. This method is further explained in (Armstrong et al. 2014) The resulting polynomial fit was then applied to the central 4hr windows, thus avoiding artificially reducing transit depths.

A search for transit signals was then performed on each lightcurve. Least-square minimisation was used to fit pre-generated transit models (developed from the Mandel & Agol 2002 small planet, quadratic limb-darkened model) to a window of the lightcurve 6 times the transit duration ($T_D$). This was repeated for transit models with durations from 1.5 to 24hrs in increasing $T_D$ steps of 25%, with each fitting window shifted by 25% of the targeted $T_D$ each time. Well-fitting models with depths greater than 2.5$\sigma$ from the out-of-transit RMS were recorded. A combination of highest SNR and lowest reduced $\chi^2$ value was used to select the best transit fit when multiple durations and transit centres were flagged on the same region of lightcurve.

To reduce false positives from thruster firings, the SNR of detections with above-average numbers of events occurring concurrently were suppressed. Lightcurves were then

\(^1\) Detrended lightcurves are publicly available on MAST at https://archive.stsci.edu/prepds/k2varcat/
adapt the transit fitting regimes of Ian Crossfield\textsuperscript{3} and the Monte Carlo Markov chain implementation \texttt{emcee} (Foreman-Mackey et al. 2013), to estimate posterior probability distributions for each monotransit signal. Quadratic limb darkening parameters adapted for the \textit{Kepler} bandpass were interpolated from stellar temperature (Sing 2010). A Gaussian prior distribution was applied to these limb darkening parameters with values and errors set from the temperature probability distribution. Errors were set to the RMS of the 4 different colour estimates to a minimum of 150K.

If eccentricity is assumed to be zero, a circular planetary period ($P_{\text{circ}}$) can be estimated from the scaled transit velocity ($v'$) and stellar density ($\rho_*$) using Kepler’s laws:

$$P_{\text{circ}} = \frac{8\pi^2 G \rho_*}{3 v'^3} = 2\pi g R_* v'^3$$

(2)

Longer period orbits (and hence lower-velocity fits) are probabilistically less likely due to transit probability ($p_t \propto \frac{2}{a} \approx \frac{a}{P_*}$), as co-planar orbits are favoured. Hence, in these cases, the forcing of fits to shorter orbits by a linear prior may not be valid. However, the increase in transit probability of a planet at distance $x$ given transiting exoplanets on orbits $g, \pi, z$, etc is a complex problem beyond the scope of this work. The non-detection of subsequent transits in the lightcurve can also be used to set a lower limit on the orbital period, and hence an upper limit on velocity. We do not apply this technique in order to produce transit fits fully independent of stellar parameters (e.g. density).

The impact parameter was limited to the range -1.2 to 1.2 to avoid walkers building up at $b = 0$. Planet-to-star radius ratios were limited to 0.25, as above this the assumption that the transiting object is fully opaque and covering a uniform region of stellar surface breaks down.

Velocity, just like the transit duration scaled to period, is directly linked to stellar density (Seager & Mallén-Ornelas 2003). For the best constrained transit models, stellar density is likely to prove the largest uncertainty. Characterising the star, therefore, is key to estimate orbital period. Such characterisation is best performed with asteroseismology, or less accurately, with spectral fitting.

Large radii planets which spend longer crossing the rim of the stellar disc are most suitable to \texttt{Namaсте} fitting as the impact parameter can be more easily distinguished. Price et al. (2015) showed that for smaller and lower signal-to-noise planet transits the uncertainty on impact parameter increases linearly, causing poor determination of perpendicular velocity and therefore eccentricity. For less well-defined fits, parameters such as impact parameter, planet-to-star ratio and velocity become correlated, as can be seen in their posterior distributions of figure 3. This is especially true for eclipses that cannot be constrained to an impact parameter less than 1.0. In these cases the fit cannot distinguish between planetary, high-velocity disc-crossing transits and lower-velocity grazing eclipses. As impact parameter increases beyond 1.0, velocity stabilises to a minimum value

\textsuperscript{2} Publicly available at \url{https://github.com/hposborn/namaste}

\textsuperscript{3} Accessed from \url{http://www.lpl.arizona.edu/ ianc/python/}
determined by transit duration. However, even for correlated and non-gaussian parameters, Namaste allows us to put probabilistic constraints on the transit fit.

An example of the resulting posterior distributions between parameters (along with model fit and residuals) can be seen in figure 3.

2.3.1 Eccentricity

For exoplanets on non-circular orbits, the circular velocity estimated by Namaste \( v_{\text{circ}} \) depends strongly on eccentricity and the argument of periastron (equation 3, Barnes 2007).

\[
V_\theta = V_{\text{circ}} \frac{1 + e \cos \theta}{\sqrt{1 - e^2}}
\]  

(3)

As the solid angle swept out by an eccentric planet’s shadow is greater than that from an equivalent circular orbit, eccentric planets are also more likely to transit. Transit probability is especially raised near periastron, suggesting circular period estimates will on average underestimate the true period. Kipping (2013) used RV planets to study the distribution of exoplanet eccentricities and showed that close-in planets (defined as \( P < 382 \) d) have a more circular distribution than long-period planets. We use these distributions to study eccentricity’s effect on transit velocity (Figure 4). Confidence intervals. We find that eccentricity increases the median velocity and its \( 1-\sigma \) confidence intervals by \( 1^{+35}_{-9}\% \) in the short case, and by \( 3^{+35}_{-9}\% \) in the general case.

Van Eylen et al. (2014) studied the discrepancy between Kepler planet densities determined by asteroseismology and those found from transit duration to determine their eccentricities. This method is analogous to the comparison of true period with those from single transit fits, and suggests Namaste could be useful for determining the eccentricities of polytransiting planets.

In this study, we note that eccentricity can add significant uncertainty to our results, but limit ourselves to estimations of circular periods \( (P_{\text{circ}}) \) which are good approximations for the majority of cases. For the short-period regime, two-thirds of planets orbit with eccentricities less than 0.2, causing substantially lower increased uncertainties. Hence, in the majority of cases, the small increase in velocity uncertainty is negligible compared to the large uncertainties from stellar density.

2.4 Stellar Parameter Fits

Where stellar parameters from spectra are unavailable, we use photometric colours to approximate the temperatures, masses and radii of EPIC stars. The majority of K2 stars have broadband photometric measurements in both the visible (e.g. Tycho B & V bands from Høg et al. (2000)) and infra-red (e.g. 2MASS J, H, K data from Cutri et al. (2003)). From the four independent colours derived from these magnitudes, stellar temperatures can be estimated (Fitzgerald 1970). This relationship assumes that no bright companions are present, which can only be confirmed with more advanced follow-up. From this temperature, and making the assumption that the star is on the main sequence, stellar models allow a stellar radius & mass to be estimated (Torres et al. 2010). Temperature uncertainty was estimated from the RMS of the temperature from each colour estimate, or 150K if the RMS was smaller than this value. Mass and radius were taken from these temperature uncertainties down to a minimum relative error of 10%.

3 RESULTS

3.1 Application to Known Kepler Systems

To test the fitting of Namaste, we applied it to single transits from the lightcurves of 6 long-period Kepler planets and KOI...
candidates. Figure 5 shows the Namaste model fits to these transits and Table 3.1 gives the output parameters compared to their published values.
**Figure 5.** Six long-period *Kepler* planets with *Namaste* fits. Best-fit models are in black, while $1 - \sigma$ error regions are in blue. x-axis is scaled to 8 transit durations, whereas the y-axis is unconstrained.

### 3.2 Application to K2 Single Transit Candidates

We applied *Namaste* to seven potential single transit events found in K2. Fits are shown in Figure 6, fit parameters in table 3.2 and full posterior distributions in Appendix A2.
Table 1. Nanast results for 4 Kepler planets and 2 KOI candidates, compared with published parameters from traditional analyses. See Appendix A1 for posterior distributions.

| EPIC203811200 | EPIC201802470 | EPIC204634789 | EPIC201631267 | EPIC203914123 | EPIC201720401 | EPIC201635312 |
|----------------|----------------|----------------|----------------|----------------|----------------|----------------|
| KepMass        | 11.896         | 11.587         | 12.292         | 12.78          | 9.039          | 14.74          | 15.143         |
| Teff(d)        | 2121.02±0.002  | 2064.11±0.005  | 2088.616±0.001 | 1996.674±0.001 | 2113.278±0.002 | 1983.114±0.002 | 1993.985±0.001 |
| b              | 0.61±0.13      | 1.07±0.48      | 0.74±0.25      | 1.05±0.09      | 0.49±0.02      | 0.12±0.01      | 0.95±0.05      |
| (Rd), Rkep     | 3.32±0.03      | 1.97±0.2      | 21.7±8.1      | 3.72±0.05      | 0.679±0.006    | 1.51±0.04      | 7.5±0.2        |
| Rp/Rs intra     | 0.56±0.002     | 0.121±0.007    | 0.085±0.008   | 0.176±0.007    | 0.149±0.001   | 0.133±0.001   | 0.24±0.038     |
| F               | 0.99994±1.5×10⁻⁶ | 0.99998±1×10⁻⁵ | 0.99994±5×10⁻⁴ | 0.99985±1.3×10⁻⁵ | 0.99985±1.7×10⁻⁴ | 1.00000±7.3×10⁻⁵ | 0.99999±9.8×10⁻⁵ |
| u1              | 0.191±0.006    | 0.175±0.01    | 0.117±0.002   | 0.099±0.005   | 0.036±0.001   | 0.119±0.014  | 0.028±0.027    |
| u2              | 0.51±0.05      | 1.09±0.05     | 0.79±0.05     | 1.45±0.41     | 1.07±0.1     | 1.14±0.1     | 2.03±0.43     |
| R(Jup)         | 0.005±0.002    | 0.027±0.002   | 0.08±0.004    | 0.10±0.004    | 0.10±0.004    | 0.10±0.004    | 0.18±0.018     |
| s.a.m.(AU)      | 1.27±0.009     | 3.71±0.04     | 0.06±0.01     | 1.03±0.2      | 3.17±0.7     | 6.17±1.1     | 3.17±0.76     |
| T2O(K)         | 5200±200       | 5100±200      | 4800±400      | 4600±200      | 4600±400      | 4800±200      | 3400±300      |
| Rstar(nc)      | 0.94±0.09      | 0.93±0.09     | 0.87±0.09     | 0.85±0.08     | 0.74±0.07     | 0.88±0.09     | 0.89±0.36     |
| Mstar(nc)      | 0.91±0.09      | 0.88±0.09     | 0.78±0.08     | 0.74±0.07     | 0.57±0.1      | 0.78±0.08     | 0.43±0.04     |

Table 2. Output median parameters for Nanast runs of 7 K2 single transit candidates, with 16% and 84% percentile bounds. See Appendix A2 for posterior distributions.
Figure 6. Six new K2 planet-like signals with Namaste fits. Best-fit models are in black with 1 − σ error regions in light blue. x-axis is scaled to 8 transit durations, whereas the y-axis is unconstrained.
4 DISCUSSION

4.1 Known Kepler Planets & Candidates

- **Kepler-51 d** -
  Planetary velocity, planet-to-star ratio and impact parameter are all well-constrained by this fit. Well-derived stellar density from Masuda (2014) also gives period errors on the order of only 10%, and within 1σ of the true period of 130.2d. The small discrepancy between true and estimated periods is likely from the overestimation of impact parameter ($b_{\text{true}} = 0.18; b_{\text{true}} = 0.094$). This further suggests Kepler-51 d is on a circular orbit.

- **Kepler-117 c** -
  Namaste fits show good agreement to the published parameters, with all parameters within errorbars. Well-constrained stellar parameters from Bruno et al. (2015) produce relatively small errors on the estimated period of 53d, which is within 5% of the true period of 50.7d.

- **Kepler-111 c** -
  This multiplanet system, validated by (Rowe et al. 2014), has less well-constrained parameters than the other Kepler test cases, with a 20% density error. This gives wide errors on the estimated period of $240^{+130}_{-130}$, although the true period of 224.8d sits well within this distribution, and suggests density uncertainties may be over-estimated.

- **Kepler-79 d** -
  An ultra-low density gas giant, Kepler-79 is a well-constrained system with masses from TTVs (Jontof-Hutter et al. 2014). In this case, impact parameter is inaccurately fitted, leading to an overestimated period greater than 1σ from the true value. The non-zero eccentricity suggested for this planet may also cause increased discrepancy.

- **KOI976.01** -
  Despite being on the Kepler candidate list, KOI976 is an eclipsing binary with a much larger radius than would be expected from an exoplanet ($8.5^+_7_{-6}$R$_{\text{Jup}}$). However, this is a good proof of concept for this technique on low-mass eclipsing stellar (or brown dwarf) companions, for which the period estimates are still valid. The long duration of the eclipse gives extremely well-fitting parameters, with output parameters such as velocity accurate to one part in 1000. However these disagree with the published KOI catalogue values in many cases, especially impact parameter, which may be due to the eclipsing body being self-luminous. This leads to an over-estimated velocity and underestimated period of $25^{+17}_{-10}$ days, although the true period is well within the errorbars.

- **KOI1431.01** -
  This candidate shows the most gaussian and closest period approximation of all the fits, with a median period within 2% of the true value. However, poor and likely over-estimated stellar density constraints give large errorbars of $\sim \pm 100d$ and stellar radii, a circular transit period was estimated and compared to the known period. In the case of eccentric planets, as was detected by (Van Eylen et al. 2014), the period estimate is likely to be significantly offset from the circular period.

Intriguingly, stellar densities derived from stellar surface gravities for a wider range (660) of Kepler planets showed a correlation between estimated period and impact parameter for grazing transits (b > 0.6). This effect is likely due to overestimation of the duration and impact parameter of the shortest transits, for example due to smeared TTVs, which would cause an increase in the estimated period. Such effects would only be present for phase-folded transits, hence unlikely to present problems for Namaste analysis.

To avoid this issue, we compared only the better constrained densities from $R_*$ and $M_*$ estimates, or directly from asteroseismology. Parameters from the 101 and 69 planets, respectively, were resampled 500 times from a Gaussian distribution derived from their uncertainties to produce a cumulative probability distribution of the estimated period compared to the true period can be seen in figure 7.

As can be seen, the less-precise density estimates from $R_*$ & $M_*$ give a shallower distribution of orbital period estimates. However both distributions peak at the true value, and suggest that the analytical technique is valid.

4.3 K2 Single Transit Events

4.3.1 EPIC203311200

Unfortunately the centre of this single transit, which is the most obvious feature of the second half of the K2 lightcurve, is coincident with a global position shift at MJD 2121.05 of 1.23 pixels. This extreme pixel shift was poorly fitted by the interpolated flux map, causing a 5-point bump in the data during the transit (see Figure 8). To fix the poor detrending for these points, we extrapolated the well-fitted region of 3D fluxmap surface to this region of centroid shift and re-ran the detrending. This had the effect of reducing systematic noise.
across the lightcurve, including those in-transit. However, six of the largest pixel shifted points were still removed from central transit by this detrending procedure.

The Namaste fit for EPIC203311200 constrains the transit to being a disc-crossing (ie, non-grazing) eclipse, and can therefore put good constraints on the size of the eclipsing body. It estimates a size of 0.51 ± 0.05R_Jup, putting it between the sizes of Uranus and Saturn. The posterior shows a two-peaked distribution in velocity and radius ratio, which may be due to a build up of walkers at the maximum velocity threshold, or an excess of distribution at b ~ 0. We cannot currently separate the correct distribution, although transit probability favours the maximum velocity (with v' ~ 4R_⋆d − 1 giving a potentially more reasonable period of P_{circ} ~ 310d). Either period would make EPIC203311200 one of the longest-period transiting exoplanets yet discovered, with the potential to exceed Kepler-421 b (704.2d, Kipping et al. (2014)).

V-J, V-H and V-K show good agreement for a 5200 ± 200K star, however (B − V) = 0.19 appears anomalously blue. A low-resolution spectrum taken with IDS on the 2.5m Isaac Newton Telescope suggests it to be a relatively metal-poor star ([Fe/H] ~ −1.0) with a temperature around 5400 ± 300K; consistent with IR colours. The low-resolution of the spectrum did not enable a log-g estimate. If no source of dilution is present, and the star is a main sequence late-G star as suggested by photometric colours and spectra, then the source of this 3mmag eclipse signal must be planetary.

4.3.2 EPIC201892470

This shallow but V-shaped eclipse occurs only 1.2d before the end of the K2 lightcurve. This allows us to place a high lower limit (> 77d) on the orbital period. A grazing eclipse is favoured by Namaste (b=1.07). The long circular period of 2790 ± 270d estimated by Namaste is likely unphysical, and could be an indication that this is a giant rather than a main sequence star. However it should be noted that the more planetary posterior regions - those with smallest R_p/R_⋆, lowest b and hence larger v' - give the shortest and therefore more likely periods.

4.3.3 EPIC204634789

With only four in-transit datapoints, the fit is poorly constrained and is unable to distinguish between any impact parameters. The velocity posterior is again double peaked because of walker build-ups at maximum velocity (b = 0) and minimum velocity where the distribution goes flat (b > 1.0). Either hypothesis gives extremely fast velocities, and therefore circular periods far shorter than can be constrained by the non-detection of other events in the lightcurve. Hence, an eclipse at the perihelion of an eccentric object (either a planet or grazing EB) is more likely. Alternatively, the signal could simply be unexplained red noise or even an outburst of evaporation from an otherwise undetectable planet.

4.3.4 EPIC201631367

Namaste suggests this shallow but V-shaped transit is most likely a grazing eclipse, although the distributions build up at the R_p/R_⋆ and b limits, suggesting the true eclipsing body could be larger than our model can fit. The period of 440 ± 160d is well-constrained, however, as the velocity of the eclipsing body becomes constant for extremely grazing eclipses.

4.3.5 EPIC203914123

The deepest and longest eclipse detected in K2, this star undergoes nearly 3 days of 2.5% dimming. Namaste, using a main-sequence approximation from photometric colours, estimates a period of 84,000d (230yr). The length of the eclipse means fitting was extremely precise and the posterior distributions are extremely well-constrained, although some red noise (seen in the form of a V-shape during the base of the transit) remains. However, the likelihood of seeing such an eclipse around a main sequence star is around 10^−7. Hence, our assumptions are likely wrong, and this is probably a low-mass star eclipsing a sub-giant or giant star. Indeed, this hypothesis also explains the ‘ramp’ seen up to and away from the eclipse as gravity darkening or ellipsoidal variation in a tidally locked binary system. As period is directly proportional to stellar density, a main sequence star can achieve better density (and therefore period) constraints for giant star eclipses such as this. Such discoveries are important finds on their own, allowing models of stellar evolution to be tested (Gaulme et al. 2013).

4.3.6 EPIC201720401

Similar to EPIC203914123, Namaste fits this eclipse extremely well but gives an unfeasibly long period for EPIC201720401 of 6300 ± 1500d. Hence a giant/M-dwarf eclipsing binary is the more likely scenario, but we cannot rule out either scenario without further follow-up, such
as radial velocities. With good estimates of stellar density, however, the period could be estimated accurately and subsequent eclipses followed-up.

4.3.7 EPIC201635132

Similarly deep to the previous eclipses, this eclipse is also short and V-shaped, which Namaste identifies as a grazing eclipse with minimum impact parameter of 0.85 and minimum radius of 0.17R\textsubscript{e}, constrained only by our imposed limits on b and R\textsubscript{p}/R\textsubscript{e}. With a minimum period of 63 days from the lightcurve, the velocity is likely to be around half of the expected value, suggesting a much more grazing eclipse than possible to detect with Namaste.

4.4 Source of Uncertainty

For our best-constrained K2 candidates, the dominant source of error is stellar density. Even for a main sequence star with a well-constrained stellar temperature, density uncertainties can be on the order of 50%. For less well-constrained probability distributions, for example from noisy lightcurves, the uncertainty in planetary velocity can be important. This is especially true for small planets for which impact parameter becomes more degenerate.

As shown in section 2.4, eccentricity also has a role in increasing period uncertainties. However, the majority of exoplanets are expected to be on circular or near-circular orbits for which the correction is minimal (Kipping 2013). If density can be constrained by follow-up observations, however, eccentricity could become an important factor when searching for a subsequent transit.

One interesting exception may be known multiplanet systems, which improve the single transit method in three ways: The low mutual inclination in such systems increases the transit probability of exterior planets; stability constraints and formation pathways also limit long-period planets to more circular orbits; and stellar parameters such as density are significantly improved due to the information gathered by interior planets (for example densities through transit durations). As such, searches for single transits in known multi-planet systems could have valuable results.

4.5 Potential for Follow-up Observations

Single spectra of candidates could immediately rule out large binary companions from cross-correlation functions. Spectral fitting would also allow better estimates of stellar parameters, including density from log-g. This would also rule out giant stars from the analysis. Lucky or AO imaging could rule out close companions that could either be false positives, or diluting the transit depth.

Estimated planetary radius and orbital period can give likely RV amplitudes and determine the observability of individual targets. Radial Velocities can be tailored to the likely period to give best observing strategy. For objects with poorly-constrained parameters (e.g. grazing transits), RVs could be used to reduce the uncertainty on planet-to-star radius ratios. This could then be used as a prior to re-calculate the planetary velocity, and hence improve the period estimate.

Good period estimates from single transit candidates could then be used to search for repeat transits. Such analysis would ideally utilise observers at different longitudes to cover a likely transit range to be many days. Long period (and low impact parameter) planets may have a better probability of re-observation due to their longer transit duration. Precise transit survey telescopes such as MEarth (Irwin et al. 2008) or NGTS (Wheatley et al. 2013), in combination with amateur observer programs, could be used most effectively to this goal.

For the smallest transit depths, small space-based telescopes such as MOST (Rucinski et al. 2003) or Cheops (Broeg et al. 2013) may be the only method of confirming the orbital period.

4.6 Validation

We have shown that, for favourable single transits, the orbit and size of a transiting exoplanet can be accurately determined, especially for giant planets. However, alternative sources for such transits (e.g. from background eclipsing binaries) remain. As has been shown (Morton 2012), the source of such false positives can be probabilistically excluded with follow-up data. For example sub-arcsecond imaging can rule out stellar companions that could be producing spurious signals. Spectral fitting to high-SNR spectra can be used to rule out closer binary companions. Radial velocity measurements (even where the detection of the signal from planet is impossible) could be used to place limits on the size of any companion. In those cases without close companions, the planetary radius of the eclipsing object can be probabilistically limited to a planetary, rather than stellar, origin. This could, in a similar way to the validation of other single transiting systems such as Kepler-452 (Jenkins et al. 2015), constitute a probabilistic validation of the planet without observing subsequent transits.

4.7 Application to Future Missions

Single transits could be used by many future surveys to detect long-period transiting planets around bright stars.

The Transiting Exoplanet Survey Satellite (TESS, Ricker et al. (2015)) will monitor 200,000 stars on 2min cadence over 28d observing windows. Studies of TESS planet yield suggest more than 100 single transits could be detected above a noise threshold of 7.3σ (Sullivan et al. 2015). Namaste could be an important tool in the follow-up of these planets.

Initial observing plans for the PLAnetary Transits and Oscillation of Stars (PLATO) mission suggest 6 fields could be observed on 2-5 month campaigns, yielding 60,000 bright stars (V mag < 12) with 30s cadence and hundreds of thousands of fainter stars (12 < V mag < 16) with 10-minute cadence. The potential combination of asteroseismology-derived densities (accurate to 10%) with high-cadence, high-precision photometric data could produce dozens of validated long-period planets (Osborn et al. prep).

Ground-based surveys such as MEarth (Irwin et al. 2008) already search for the single transits of stars with
an innovative detection strategy (Nutzman & Charbonneau 2008). This method could also be applied to more traditional single-field photometric surveys such as NGTS (Wheatley et al. 2013).

Gaia has the capability to detect tens of thousands of giant planets on 1 to 4 year orbits (Dzigan & Zucker 2012), hence any transiting gas giants on these orbits found by TESS, PLATO or from the ground stand a good chance of being confirmed by Gaia astrometry. Hence this method could lead to the first overlap between the realms of transiting and astrometric exoplanet astronomy.

The long transit duration of planets detected from a single transit also makes them amenable to transmission spectroscopy, for example by JWST (Belu et al. 2011). With a dearth of transiting warm and cold Jupiters known around bright stars, such objects could prove a vital link between the atmospheres of hot Jupiters and those of solar System gas giants.

5 CONCLUSION

We have developed Namaste, a method of combining stellar parameters with the lightcurve of a single transit to estimate orbital parameters. We have tested this analysis on published transit parameters for a large sample of Kepler multiplanets, showing close agreement. A test of the full fit-published transit parameters for a large sample of Kepler mate orbital parameters. We have tested this analysis on parameters with the lightcurve of a single transit to esti-

Namaste is therefore a useful tool, allowing more targeted follow-up observations with radial velocity measurements and the search for additional transits. For giant planets around bright stars, subsequent RV follow-up could lead to their confirmation. Smaller planets around fainter stars could still be validated with follow-up techniques such as diffraction-limited imaging and high-resolution spectroscopy.

In future the detection and analysis of single transits in this way could lead K2, TESS, PLATO and ground-based photometric surveys to detect transiting exoplanets on orbits longer than their observing campaigns would traditionally allow. Such long-period planets, especially if found around bright stars, could pave the way for a new regime of exoplanetary science. This includes the detection of planets within the astrometric sensitivity of Gaia, and cold Jupiters with atmospheres observable in transmission spectroscopy by JWST. EPIC203311200 b, if proved to be planetary by ongoing follow-up work, could fit both roles.

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APPENDIX A: MCMC POSTERIOR DISTRIBUTIONS

This paper has been typeset from a TeX/LaTeX file prepared by the author.
Figure A4. Posterior distribution for Kepler-79 d (KIC8394721).

Figure A5. Posterior distribution for KOI 976.01 (KIC34417842).

Figure A6. Posterior distribution for KOI 1431.01 (KIC11075279).

Figure A7. Posterior distribution for EPIC203311200.
Figure A8. Posterior distribution for EPIC201892470).

Figure A9. Posterior distribution for EPIC204634789).

Figure A10. Posterior distribution for EPIC201631267).

Figure A11. Posterior distribution for EPIC203914123).
Figure A12. Posterior distribution for EPIC201720401).

Figure A13. Posterior distribution for EPIC201635132).