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The **XMM-Newton** serendipitous survey

**IX. The fourth XMM-Newton serendipitous source catalogue**

N. A. Webb, M. Coriat, I. Traulsen, J. Ballet, C. Motch, F. J. Carreras, F. Koliopanos, J. Authier, I. de la Calle, M. T. Ceballos, E. Colomo, D. Chuard, M. Freyberg, T. Garcia, M. Kolehmainen, G. Lamer, D. Lin, P. Maggi, L. Michel, C. G. Page, M. J. Page, J. V. Perea-Calderon, F.-X. Pineau, P. Rodriguez, S. R. Rosen, M. Santos Lleo, R. D. Saxton, A. Schwepe, L. Tomás, M. G. Watson, and A. Zakardjian

1. Introduction

The sky is constantly being surveyed by many different telescopes exploiting the full range of the electromagnetic spectrum, in addition to gravitational wave, neutrino, and cosmic ray observatories. Each observation can provide a clue as to the nature of the source and the physical processes underway. In addition, many objects are known to be highly variable in time, requiring many observations to fully understand the nature of the variability. Whilst dedicated observations can be necessary to answer some science questions, frequently, catalogues can provide the required information. Catalogues can also provide homogeneous datasets for classes of objects as well as reveal previously unknown objects.

Catalogues have been produced for the majority of the X-ray missions that have flown. Early X-ray missions detected very few objects. The fourth version of the UHURU catalogue (1970–1973, Forman et al. 1978) indicates that just 339 X-ray sources were discovered by the satellite. The HEAO 1 catalogue (1977–1978, Wood et al. 1984) provides 842 X-ray sources. The ROSAT catalogue, 2RXS (1990–1991, Boller et al. 2016) gives 135,000 X-ray detections or 129,192 sources. However, more recent X-ray observatories have several advantages over the earlier missions. Firstly, they have a larger collecting area
and are therefore more sensitive. Secondly, they have also surveyed the sky for a much longer period and thus they detect many more sources. *Chandra*, which was launched in July 1999, boasts a very extensive catalogue, the *Chandra* Source Catalog Release 2.0 (CSC 2.0; Evans et al. 2014; Chen et al. 2019; Evans & Civano 2018; Wilkes et al. 2019) with 928 280 X-ray detections, which are from 317 167 individual X-ray sources. The *Neil Gehrels Swift* Observatory was launched in November 2004 and the 2SXPS catalogue (Evans et al. 2020) lists 1.1 million detections, which are of 206 335 individual X-ray sources. The major advantage of this catalogue is that it covers a larger lion of sky, and sources have been pointed many times over the last 16 years. The first catalogue of sources detected with the hard X-ray observatory, *NuSTAR*, lists 497 sources after 40 months of observations (Lansbury et al. 2017).

This paper focuses on the catalogue of detections from the European Space Agency’s second cornerstone mission from the Horizon 2000 programme, *XMM-Newton* (Jansen et al. 2001), which was launched twenty years ago on 10 December 1999. It has the largest effective area of any X-ray satellite (Ebrero 2019) thanks to the three X-ray telescopes aboard, each with $\sim$1500 cm$^2$ of geometric effective area. This fact, coupled with the large field of view (FOV) of 30’ diameter, means that a single pointing with the mean duration in the catalogue of 37 ks detects 70–75 serendipitous X-ray sources. The catalogue of serendipitous sources from overlapping *XMM-Newton* observations 4XMM-DR9s is described in paper X of this series (Traulsen et al. 2020).

The *XMM-Newton* Survey Science Centre\(^1\) (SSC), a consortium of ten European Institutes (Watson et al. 2001), has developed much of the *XMM-Newton* Science Analysis System (SAS; Gabriel et al. 2004) used for reducing and analysing *XMM-Newton* data and created pipelines to perform standardised routine processing of the *XMM-Newton* science data. The XMM-SSC also produces catalogues of all of the detections made with *XMM-Newton*. The catalogues of X-ray detections made with the three EPIC (Strüder et al. 2001; Turner et al. 2001) cameras that are placed at the focal point of the three X-ray telescopes have been designated 1XMM, 2XMM, and 3XMM (Watson et al. 2009), with incremental versions of these catalogues indicated by successive data releases, denoted-DR in association with the catalogue number. This paper presents the latest version of the XMM catalogue, 4XMM, which spans 19 years of observations made with *XMM-Newton* and includes many improvements with respect to previous *XMM-Newton* catalogues. The most notable change between 3XMM and 4XMM is the methodology used for background modelling (see Sect. 3.4).

### Table 1. Characteristics of the 11 204 *XMM-Newton* observations included in the 4XMM-DR9 catalogue.

| Camera | Modes | Filters | Total |
|--------|-------|---------|-------|
|        | Full \(^a\) Window \(^b\) Other \(^c\) | Thin | Medium | Thick |
| pn     | 8462 | 683 | 1344 | 5640 | 4011 | 838 | 10 489 |
| MOS1   | 8681 | 1950 | 373 | 5080 | 4943 | 981 | 11 004 |
| MOS2   | 8728 | 1981 | 348 | 5120 | 4971 | 966 | 11 057 |

**Notes.** \(^a\)Prime Full Window Extended (PFWE) and Prime Full Window (PFW) modes; \(^b\)pn Prime Large Window (PLW) mode and any of the various MOS Prime Partial Window (PPW) modes; \(^c\)other pn modes such as the Small Window, timing or burst modes, MOS modes (Fast Uncompressed (FU), Refresh Frame Store (RFS)).

**Fig. 1.** Hammer-Aitoff equal area projection in Galactic coordinates of the 11 204 4XMM-DR9 fields.

3. Data processing

Data processing for the 4XMM-DR9 catalogue was based on the SAS version 18 and carried out with the pipeline version 18\(^2\) and the latest set of current calibration files at the time of processing (February and March 2019).

The main data processing steps used to produce the 4XMM data products were similar to those outlined in Rosen et al. (2016), Watson et al. (2009) and described on the SOC webpages\(^3\). For all the 4XMM data, the observation data files were processed to produce calibrated event lists. The optimised background time intervals were identified and using them, the filtered exposures (taking into account exposure time, instrument mode, etc.), multi-energy-band X-ray images, and exposure maps were generated. The source detection was done simultaneously on all

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\(^1\) [http://xmmssc.irap.omp.eu/](http://xmmssc.irap.omp.eu/)

\(^2\) [https://www.cosmos.esa.int/web/xmm-newton/pipeline-configurations](https://www.cosmos.esa.int/web/xmm-newton/pipeline-configurations)

\(^3\) [https://xmm-tools.cosmos.esa.int/external/xmm_products/pipeline/doc/17.40_20181123_1545/modules/index.html](https://xmm-tools.cosmos.esa.int/external/xmm_products/pipeline/doc/17.40_20181123_1545/modules/index.html)
images and bands, one to five, from the three cameras as in Watson et al. (2009), Rosen et al. (2016). The probability, and corresponding likelihood, were computed from the null hypothesis that the measured counts in the search box result from a Poissonian fluctuation in the estimated background level. A detection mask was made for each camera that defines the area of the detector which is suitable for source detection. An initial source list was made using a “box detection” algorithm. This slides a search box (20′′ × 20′′) across the image defined by the detection mask. Sources were cut-out using a radius that was dependent on source brightness in each band, and these areas of the image where sources had been detected were blanked out. The source-excised images, normalised by the exposure maps, and the corresponding masks are convolved with a Gaussian kernel to create the background map (see Traulsen et al. 2019, where this smoothing method is new for the detection catalogue). A second box-source-detection pass was then carried out, creating a new source list, this time using the background maps (“map mode”) which increased the source detection sensitivity compared to the first pass. The box size was again set to 20′′ × 20′′. A maximum likelihood fitting procedure was then applied to the sources to calculate source parameters in each input image, by fitting a model to the distribution of counts over a circular area of radius 60′′.

Table 2. Data modes of XMM-Newton exposures included in the 4XMM catalogue.

| Abbr. | Designation | Description |
|-------|-------------|-------------|
| MOS cameras: | | |
| PFW | Prime Full Window | Covering full FOV |
| FPW2 | Prime Partial W2 | Small central window |
| FPW3 | Prime Partial W3 | Large central window |
| FPW4 | Prime Partial W4 | Small central window |
| FPW5 | Prime Partial W5 | Large central window |
| FU | Fast Uncompressed | Central CCD in timing mode |
| RFS | Prime Partial RFS | Central CCD with different frame time (“Refreshed Frame Store”) |
| pn camera: | | |
| PFWE | Prime Full Window Extended | | |
| PFW | Prime Full Window | Covering full FOV |
| PLW | Prime Large Window | Half the height of PFW/PFWE |

3.1. Exposure selection

The same criteria used for selecting exposures for 3XMM were retained for 4XMM. A total exposure time of 410 Ms was available for 4XMM-DR9, with an increase of 57% compared to 3XMM-DR5.

3.2. Event list processing

Much of the pipeline processing that converts raw ODF event file data from the EPIC instruments into cleaned event lists has remained unchanged from the pre-cat9.0 pipeline and is described in Sect. 4.2 of Watson et al. (2009). A number of improvements have been made since the 2XMM (Watson et al. 2009) and 3XMM (Rosen et al. 2016) catalogues, which can be found in the SAS release notes 5. These include source spectra and light curves created for pn Timing mode and small window data, source detection on pn small window data, energy dependent Charge Transfer Inefficiencies (CTI) and double event energy corrections, time and pattern dependent corrections of the spectral energy resolution of pn data, X-ray loading and rate dependent energy (PHA) and CTI corrections for EPIC pn Timing and Burst modes, binning of MOS spectra changed from 15 eV to 5 eV. Filtering was carried out with XMMEA_EM, which is a bit-wise selection expression, automatically removing bad events such as bad rows, edge effects, spoiled frames, cosmic ray events (MIPs), diagonal events, event beyond threshold, etc, instead of XMMEA_SM (which removed all flagged events except those flagged only as CLOSE_TO_DEADPIX). Other modifications include the generation of background regions for EPIC spectra and light curves selected from the same EPIC chip where the source is found, observations of solar system objects processed such that X-ray images and spectra correctly refer to the moving target, the inclusion of pileup diagnostic numbers for EPIC sources (see also Sect. 6.4.1), and footprints for EPIC observations based on combined EPIC exposure maps provided as ds9 region files. Other changes carried out specifically for the production of 4XMM include a revised systematic position error (see Sect. 3.3), the modelling of the EPIC background (see Sect. 3.4) and finer binning of EPIC lightcurves (see Sect. 4.1).

4 https://nxsa.esac.esa.int/nxsa-web

5 https://www.cosmos.esa.int/web/xmm-newton/sas-release-notes/
A small rotation of ~0.4° was noted in 3XMM fields, but analysis of 4XMM data shows that the recent improvements to calibration have resolved this issue. Below we describe some of the more recent developments specifically implemented for 4XMM.

3.3. Systematic position error

The astrometry of the X-ray detections is improved by using the catcorr task to cross-correlate the X-ray detections with the USNO B1.0, 2MASS or SDSS (DR8) optical or IR catalogues. Using pairs of X-ray and optical or infra-red detections that fall within 10″ of each other, the astrometry for the field is corrected using a translational shift in the right ascension (RA) and declination (Dec) directions, together with the rotational error component. A systematic error on the position (SYSERRCC) is then calculated using the 1σ errors on the shifts in the RA (Δα error) and Dec (Δδ error) directions and the rotational error component in radians (Δθ error), derived from from the catalogue that yields the “best” solution, using

\[ \text{SYSERRCC} = \sqrt{\text{Δα error}^2 + \text{Δδ error}^2 + (r + \text{Δθ error} F)} \]

where \( r \) is the radial off-axis angle of the detection from the spacecraft boresight in arcsecs. However, where catcorr fails to obtain a statistically reliable result (poscorrok=false), a systematic error of 1.5″ was used to create the 3XMM catalogue.

In the framework of creating 4XMM, this systematic error was re-evaluated. In order to determine an improved systematic error, we identified fields in 3XMM-DR8 where catcorr failed. We used sources from the Sloan Digital Sky Survey Data Release 12 quasar (SDSS DR12 QSO) catalogue (Pâris et al. 2017) with good quality spectra (ZWARNING=0) and point-like morphology (SDSS MOPHBO=0). To avoid mismatches between targets and matched photometry we chose non-empty OBJ_ID values. We then cross-matched with the SDSS DR9 photometry catalogue (Ahn et al. 2012) in Vizier7 with a maximum distance of 5″. This step provided the uncertainty in the astrometric position of SDSS. We adopted the radially-averaged uncertainty in the SDSS positions to which we had already added a systematic 0.1″ in quadrature, \( \Delta \alpha = \sqrt{(\Delta \alpha_X^2 + \Delta \alpha_Y^2)/2 + 0.1^2} \). We then discarded all quasars with more than one SDSS DR9 counterpart within 5″. Out of the 256107 “clean” quasars, we selected the potential counterparts to the 3XMM DR8 sources, but also discarded those which could be counterparts of more than one 3XMM DR8 source. We used the “slim” catalogue for this purpose, since multiple detections of the same physical source appear only once. The combined positional error on each source in the slim catalogue is \( \Sigma = \text{SC\_POSSERR} \), calculated as the weighted average of the individual positional errors \( \text{POSSERR} \) of the individual detections. In turn, this is calculated as \( \text{POSSERR} = \sqrt{\text{RADEC\_ERR}^2 + \text{SYSERRCC}^2} \), where

\[ \text{RADEC\_ERR} = \sqrt{\Delta \alpha_X^2 + \Delta \delta_X^2} \] (\( \Delta \alpha_X \) and \( \Delta \delta_X \) are the 1σ uncertainties in the RA and Dec coordinates, respectively).

We cross-matched the SDSS DR9 positions of “clean” QSOs with the positions of the sources in the slim catalogue out to a distance of \( r = 30′′ \). For each of the resulting pairs we estimated the combined positional error as \( \sigma = \sqrt{\Delta \alpha^2 + \Delta \delta^2/2} \), where \( \Delta \alpha = \text{SC\_POSSERR} \) and discarded all quasars that had more than one counterpart out to \( r = 6′′ \), leaving 7205 suitable QSO (there were 26 QSO with more than one counterpart out to that limit). There were no pairs of quasars that corresponded to the same X-ray source.

Since each instance of an X-ray source in the 3XMM-DR8 detection catalogue is an independent measurement, we cross-matched the sample of suitable quasars with the detection catalogue where poscorrok=false, out to \( r = 30′′ \) again, filtering the latter with SUM\_FLAG=0 and EP\_EXTENT=0, to keep only the cleanest sample of secure point-like X-ray sources. At this point we have 178 quasar-X-ray detection pairs. As for the slim catalogue, we define the combined positional error as \( \sigma = \sqrt{\Delta \alpha^2 + \Delta \delta^2/2} \), where \( \Delta \alpha = \text{RADEC\_ERR} \) and \( x = r/\sigma \). Our final filtering retained only the 157 QSO-X-ray pairs with \( x < 5 \).

The expected probability density distribution of \( x \) should follow the Rayleigh distribution \( P(x) = xe^{-x^2} \). Since this was not the case for the 157 pairs of sources found above, we added an additional positional uncertainty, \( \Sigma \), in quadrature, so that the total positional uncertainty is now \( \sigma’ = \sqrt{\sigma^2 + \Sigma^2} \), looking for the value of \( \Sigma \) that minimises the difference between the distribution of the \( x’ = r/\sigma’ \) and the Rayleigh distribution using maximum likelihood. We found \( \Sigma = 1.29 \pm 0.12′′ \), where the uncertainty (1σ) has been calculated by bootstrap with replacement. The improvement can be seen in Fig. 3. This value was then used to replace the 1.5′′ systematic error when poscorrok=false. We note that a minor error was introduced into 4XMM-DR9, where the systematic error used in the case of poscorrok=false was a factor \( \sqrt{2} \) too small. This is corrected in versions 4XMM-DR10 and higher.

3.4. Modelling the EPIC background

For each input image to the source detection, the background is modelled by an adaptive smoothing technique. The method was initially applied to the data in the 3XMM-DR7 catalogue which treats overlapping XMM-Newton observations and is described by Traulsen et al. (2019). Since 3XMM-DR7 was based on a selection of clean observations, the smoothing parameters were revised for the 4XMM catalogues, which cover observations of all qualities. The three parameters of the smoothing task are the cut-out radius to excise sources, the minimum kernel radius of the adaptive smoothing, and the requested signal-to-noise ratio in the map. Their best values were determined in a three-fold assessment which involved real observations, randomised images, and visual screening.

656 observations were chosen which cover positions of cluster candidates identified by Takey et al. (2013) to involve a considerable number of extended and point-like sources. Their...
background was modelled using different combinations of the smoothing parameters, and source detection was performed. The number of detections and recovered clusters, and the source parameters of the clusters and point-like detections were compared, opting for a reasonable compromise between total number of detections and potentially spurious content and for reliable fluxes, and extent radius of the clusters. The source parameters of point-like detections were largely unaffected by the parameter choice in the tested parameter range.

The optimisation was then re-run on ninety observations, in which the background was replaced by a Poissonian randomisation. Finally, the two best combinations of smoothing parameters and the previously used spline fit were compared in a blind test. The detection images were inspected in randomised order, so the screeners could not know which source-detection results were based on which background model. The three parts of the assessment confirmed the preference for the adaptive smoothing approach over a spline fit and the estimation of the final parameters: a brightness threshold for the source cut-out radius of $2 \times 10^{-4}$ counts arcsec$^{-2}$, a minimum smoothing radius of ten pixels (40″ in default image binning), and a signal-to-noise ratio of 12.

3.5. Updated flagging procedures

A single change to the flags provided for each detection has been introduced. Flag 12 now indicates if the detection falls on a region of the detector that can show hot pixels that can be misinterpreted as a source. The rest of this section describes how the hot areas in the detector plane are identified and flagged.

Warm pixels on a CCD (at a few counts per exposure) are too faint to be detected as such by the automatic processing, but can either push faint detections above detection level, or create spurious detections when combined with statistical fluctuations. This is an intrinsically random process, not visible over a short period of time, but which creates hot areas when projecting all detections detected over 18 years onto the detector plane.

We addressed this by projecting for each CCD all detections onto chip coordinates (pn or M1 or M2_RAWX or Y), keeping only detections above the detection threshold with the current instrument alone. In that way, we can distinguish hot areas coming from different instruments, see Fig. 4.

We proceeded to detect hot pixels or columns in each CCD, using a similar method to the SAS task embadpixfind. Because the localisation precision of faint detections is several arcseconds (larger than the MOS pixel size of 1.1″) the detection was carried out for MOS after binning the image to 3×3 pixels (and testing all 9 single-pixel shifts). We flagged hot pixels with a probability less than $10^{-2}/N_{\text{final}}$ to be compatible with a Poisson distribution at the local average (estimated from the local median plus 1). The trials factor $N_{\text{final}}$ was set to the image size ($64 \times 200$) for pn and three times the binned image size ($3 \times 200^2$) for MOS, accounting approximately for the fact that the shifted binned images are correlated.

Hot columns are detected in the same way after projecting the images (with hot pixels masked) onto RAWX. A column was considered bright when it was too high at the 7σ level applying the likelihood ratio test for Poisson counts (Li & Ma 1983) with respect to its surroundings (excluding immediate neighbours). This very high threshold was chosen such that subtle increases not obvious by eye were not detected (there are hundreds of detections per column, so that method is very sensitive).

It often occurs that only a piece of a column is bright. In order to identify such occurrences we compared the distribution of detections along RAWY in the hot column with that in the

Fig. 4. Same part of the focal plane (lower left in pn detector coordinates) viewed by pn (CCD 11) and MOS1 (CCD 2). The maps are in CCD coordinates, but offset and zoomed so that they are approximately aligned (a given detection appears at the same place on both maps). All point-like 4XMM detections with log(likelihood) XX_8_DET_ML > 6.5 in the total band for the current instrument (XX) are accumulated on each map. The MOS1 map is smoothed with a 3×3 boxcar average. The colour scale is square root between zero and three detections per pixel in MOS1, 0 and 100 in pn. Obvious hot areas are visible. They appear in only one instrument because the detections on hot areas have DET_ML > 6.5 only in the instrument where the hot area is, contrary to real detections. MOS2 is omitted because it shows no hot area in that part of the focal plane.

same neighbouring columns used in the column detection, using the Kolmogorov-Smirnov (hereafter KS) test. If the probability of compatibility was less than $10^{-2}$, we looked for the bright interval with repeated KS tests on restricted lengths on each side of the RAWY value where the maximum distance between the two distributions occurs, until we reached a probability of compatibility larger than $10^{-2}$. The remainder on each side was considered normal or hot depending on the result of a Li & Ma test at the 3σ level with respect to the neighbouring columns.

We defined contiguous hot areas after reprojecting all the hot pixels and segments of columns onto the CCD (at the full pixel resolution for MOS). Many of those warm pixels were not present at the beginning of the mission, and some appear for a short amount of time. So we tested each hot area for variability using revolution number, and the same KS-based algorithm used to detect segments of bright columns, compared to the reference established over all detections on all CCDs and all instruments. This resulted in a revolution interval for each hot area. The distribution of remaining detections is shown in Fig. 5.

Detections on a hot area for a particular instrument and within the corresponding revolution interval are flagged with flag 12. This results in 16,503 flagged sources for pn, 6,245 for MOS1 and 1382 for MOS2.

4. Source-specific product generation

In order to minimise any contribution from soft proton flares, Good Time Interval (GTI) filtering is carried out. This is done for each exposure. A high energy light curve (from 7 keV to 15 keV for pn, >14 keV for MOS) is created, and initial background flare
GTIs are derived using the optimised approach employed in the SAS task, bgkoptrate (Rosen et al. 2016). Bkgoptrate determines the background count rate threshold at which the data below the threshold yields a maximum signal to noise ratio, by filtering the periods of time when the lightcurve count rate is above the optimised threshold. Following the identification of so called bad pixels, event cleaning and event merging from the different CCDs, an in-band (0.5–7.5 keV) image is then created, using the initial GTIs to excise background flares. After source detection, an in-band light curve is generated, excluding events from circular regions of radius 60′′ for sources with count rates <0.35 ct s\(^{-1}\) or 100′′ for sources with count rates >0.35 ct s\(^{-1}\), centred on the detected sources. The SAS task, bgkoptrate, is then applied to the light curve to find the optimum background rate cut threshold and this is subsequently used to define the final background flare GTIs. If no lightcurve can be generated, a general filtering is applied for the observation. Image data are extracted from events using GTIs determined from when the pointing direction for the observation is carried out. Image data are rebinned from events using GTIs determined from when the pointing direction is within 3′ of the nominal pointing position for the observation.

Following the source detection process, detections identified with at least 100 EPIC counts have their spectra extracted. If the number of counts not flagged as “bad” (in the sense adopted by XSPEC) is still greater than 100 counts, a spectrum and a time series are extracted using an aperture around the source whose radius is automatically determined to maximise the signal-to-noise of the source data. This is done with a curve-of-growth analysis, performed by the SAS task, regionanalyse. The algorithm then searches for a circular background region on the same CCD where the source is located, excluding regions where sources have been detected, as described in Rosen et al. (2016). The exception is in the case when the source falls on the central CCD of a MOS observation in SmallWindow mode (PrimePartialW2 or 3). In that case the background is estimated from an annulus (inner radius of 5.5′ and outer radius of 11′) centred on the source. The background is therefore estimated from the peripheral CCDs and the central CCD is completely excluded. For EPIC-pn sources, the algorithm avoids the same RAWY column as the source in order to exclude out-of-time events from the background estimation. The background region always has a radius larger than three pixels, otherwise no background is calculated. Response files (.rmf and .arf) are then created using the SAS tasks rmfgen and arfgen. The pile-up is estimated as described in Sect. 6.4.1 and written to the header.

4.1. Lightcurve generation
Lightcurves are corrected using the SAS task, epiclccorr, to take into account events lost through inefficiencies due to vignetting, bad pixels, chip gaps, PSF and quantum efficiency, dead time, GTIs, and exposure. epiclccorr also takes into account the background counts, using the background lightcurve, extracted over the identical duration as the source lightcurve. The time bin size for the pn lightcurves was previously set to a minimum of ten seconds and could be as poorly sampled as tens of thousands of seconds for the faintest sources. To exploit the high time resolution and high throughput of the pn, for 4XMM we now extract the pn lightcurve such that each bin is 20 times the frame time, usually 1.46 s. The binning of the MOS data remains as it was for 3XMM.

4.2. Variability characterisation
As in 3XMM, the \( \chi^2 \) test was used to determine if a source is variable during a single observation. Variability was defined as \( P(\chi^2) \leq 10^{-3} \). We also gave the fractional variability, \( F_{\text{var}} \), to provide the scale of the variability (Rosen et al. 2016). These values are still provided in the 4XMM catalogue. The \( \chi^2 \) statistic can be applied to binned data sets where the observed number of counts in a bin deviates from expectation approximately following a Gaussian distribution. Cash (1979) showed that when the number of counts per bin falls below \( \sim 10\)–20, the \( \chi^2 \) statistic becomes inaccurate. Therefore, as the pn lightcurves are now binned to 20× frame time (Sect. 4.1), these data are rebinned to contain 20 counts per bin before applying the variability tests. Future versions of the catalogue are expected to exploit the high time resolution of the pn lightcurves using a Kolmogorov-Smirnov test which can be carried out on finely binned data.

As in previous catalogue versions, we still provide columns with the minimum EPIC source flux (and error) and the maximum EPIC source flux (and error). This allows the user to find sources variable between observations. Alternatively, the fluxes from each observation, along with the observation date can be seen directly as a table when querying a source on the catalogue server\(^8\). Whilst the majority of sources do not vary in flux, the maximum variability in the catalogue is a factor \( \times 10^3 \) in flux (for example, V2134 Oph a low mass X-ray binary). Variability between observations is also provided in the stacked catalogue, see Sect. 7.

5. Screening
Visual inspection of each detection in every observation that was included in 4XMM was carried out, as has been done for previous versions of the catalogue (Rosen et al. 2016). The aim of the screening is to visually validate the new methodology employed in the pipeline, ensure that the pipeline processing has run correctly, and to flag detections that are likely to be spurious and that have not been automatically identified as possibly spurious in the pipeline processing. Whilst the source detection process is very robust, some spurious detections can still occur in the

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\(^8\) http://xmm-catalog.irap.omp.eu/
6. Catalogue construction

6.1. Unique sources

The 4XMM detection catalogue contains multiple detections (up to 69 times in the most extreme case) of many X-ray sources, due to partial overlap between fields of view as well as repeated observations of the same targets. As has been done in previous versions of the catalogue (Rosen et al. 2016), we assign a common unique source identifier, SRCID, to individual detections that are considered to be associated with the same X-ray source on the sky. The procedure used to perform associations is the same (and therefore subject to the same caveats) as the one outlined in Sect. 6 of Rosen et al. (2016).

6.2. Naming convention for the DETID and the SRCID

Starting in 3XMM-DR5, the procedure for attributing the detection identification number (DETID) and the unique source identification number (SRCID), both being unique to each detection and each unique source respectively, has been modified. Previously, identification numbers were re-computed for each catalogue version leading to supplementary columns added to the catalogue with the DETID and SRCID from previous releases.

The DETID is now constructed from the OBS_ID, which always remains the same for an observation, coupled with the source number SRC_NUM as follow:

\[ \text{DETID} = "1" + \text{OBS_ID} + \text{SRC_NUM} \]

where the “+” sign indicates string concatenation and where SRC_NUM is zero-padded to form a four digit number. The SRCID of a unique source is then determined from the first DETID attributed to that source (that is, in the observation where the source was first detected) and replacing the first digit “1” by “2”.

Despite the new naming convention that aims at preserving SRCID numbers across catalogue versions, a certain number of SRCID can disappear from one catalogue version to another. This is a normal consequence of the algorithm that groups detections together into unique sources (see Sect. 6 of Rosen et al. 2016). When new data are added and statistics are improved, the algorithm might find a better association of detections into unique sources. As an example, a total of 134 SRCIDs listed in 3XMM-DR7 are absent in 3XMM-DR8.

6.3. Missing detections and DETID change

In addition, the pipeline reprocessing of the full public dataset for the 4XMM version of the source catalogue led to significant modifications of the detection list. There are 10 214 observations that are common between the 3XMM-DR8 and 4XMM-DR9 catalogues, resulting in 775 241 detections in 3XMM-DR8 and 726 279 detections in 4XMM-DR9. Of these, there are 608 071 point-like detections with a SUM_FLAG < 1 in 3XMM-DR8 and 607 196 in 4XMM-DR9. However, amongst these observations, there are ~128 000 detections that appear in 3XMM-DR8 that are not matched with a detection in the same observation in the 4XMM-DR9 catalogue within a 99.73% confidence region (that is, 2.27 × POSERR). About 67 000 of these were classified as the cleanest (SUM_FLAG < 1), point-like sources in 3XMM-DR8 – these are referred to as missing 4XMM detections in what follows. Conversely, there are ~164 000 detections in the 4XMM-DR9 catalogue that are in common observations but not matched with a detection in 3XMM-DR8 within 99.73% confidence region, approximately 107 000 of which are classified as being clean and point-like. This is an expected consequence of the reprocessing which was already encountered in the transition from 2XMM to 3XMM (see Sect. 8 and Appendix D in Rosen et al. 2016). The number of missing 4XMM detections is consistent with the number of missing 3XMM detections, where there were ~25 700 good detections that appeared in 2XMMI-DR3 that were not matched with a detection in the same observation in the 3XMM-5 catalogue (Rosen et al. 2016).

This amounts to ~4.5% which is of the same order as the number of missing sources in 4XMM (8.3%). The origin of these source discrepancies between the two catalogues are the improvements made to the pipeline and in particular the new background estimation. The majority of the detections present in 3XMM-DR8 that are not present in 4XMM-DR9 are from the lowest maximum likelihoods, see Fig. 6. A small change in the parameters can cause a source with a maximum likelihood close to the cutoff of 6, but no the less slightly above, to have a value slightly below the cutoff and therefore be excluded from the catalogue. Conversely, the changes in the pipeline for sources just below the

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Table 3. 4XMM observation classification.

| OBS CLASS | Masked fraction | 3XMM-DR8 | 4XMM-DR9 |
|-----------|----------------|---------|---------|
| 0         | bad area = 0%  | 18%     | 30%     |
| 1         | 0% < bad area < 0.1% | 17% | 30%     |
| 2         | 0.1% < bad area < 1% | 16% | 17%     |
| 3         | 1% < bad area < 10% | 26% | 17%     |
| 4         | 10% < bad area < 100% | 14% | 9%      |
| 5         | bad area = 100% | 4%      | 1%      |

Notes. OBS_CLASS, is the observation class given in the first column, the percentage of the field considered problematic is in the second column and the percentage of fields that fall within each class for 3XMM-DR8 and 4XMM-DR9 are given in the third and fourth columns respectively.

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SRC_NUM is the source number in the individual source list for a given observation; Sources are numbered in decreasing order of count rate (that is, the brightest source has SRC_NUM = 1).
maximum likelihood cut-off of six and therefore not in 3XMM-DR8 can mean that they will then have a higher maximum likelihood and be present in 4XMM-DR9. As discussed in Sect. 5, fewer obviously spurious detections are found in 4XMM-DR9 than in 3XMM-DR8, which is also reflected in Fig. 6, where the detections found in 4XMM-DR9 and not in 3XMM-DR8 are generally more reliable (higher maximum likelihood).

A related consequence is that the source numbering within a given observation (that is, the SRC_NUM) has been altered in 4XMM-DR9 by the detections added and those removed. Therefore, amongst the detections that are matched between 3XMM-DR8 and 4XMM-DR9, the majority of them have different DETIDs in 4XMM-DR9 and 3XMM-DR8 (since the DETID is constructed from SRC_NUM). To minimise this effect, for the detections matched between the two catalogues, we have chosen to keep the original 3XMM-DR8 DETIDs instead of the newly generated ones for 4XMM-DR9. However, in doing so, we ended up with ~36 000 DETID duplicates due to unmatched 4XMM-DR9 detections having the same DETID as matched 3XMM-DR8 detections. In such cases, we added 5000 to the DETID of the unmatched detection to create a new unique DETID.

6.4. New and revised data columns in 4XMM

We have taken the opportunity of this major release version to revise some data columns and introduce new ones to the catalogues of detections and unique sources (the slen version). A pileup evaluation per instrument for each detection is now provided as three new columns: PN_PILEUP, M1_PILEUP, and M2_PILEUP, see Sect. 6.4.1. In 3XMM-DR8 and earlier versions, the extent likelihood EP_EXTENT_ML was provided only for sources detected as extended. We now provide the extent likelihood for all sources, see Sect. 6.4.2. The source extent of unique sources (SC_EXTENT) is now calculated using a weighted average. We now provide the error on the total band extent of a unique source: SC_EXT_ERR. It is calculated in the same way as the errors on the other unique source parameters (for example, the SC_EP_FLUX_ERR or the SC_HRn_ERR) namely, as the error on the weighted mean:

\[
SC\_EXT\_ERR = \left( \frac{1}{\sum_i \frac{1}{EP\_EXTENT\_ERR_i}} \right)^{-1/2}
\]

where EP_EXTENT_ERR is the total band error on the extent of the ith detection of the unique source.

6.4.1. Pile up information

As of 4XMM we provide three new columns (PN_PILEUP, M1_PILEUP, and M2_PILEUP) quantifying whether each detection may be affected by pile-up in any instrument. A value below unity corresponds to negligible pile-up (less than a few \% flux loss) while values larger than ten denote heavy pile-up. Pile-up is dependent on time for variable detections. We neglect that here, but we note that a variable detection is more piled-up than a constant one for the same average count rate, so our pile-up level can be viewed as a lower limit. We also neglect the slight dependence on the detection spectrum due to the event grade dependence of pile-up.

Our pile-up levels are not based on a fit of the full images using a pile-up model (Ballet 1999). For point sources, they are equal to the measured count rates reported in the catalogue over the full energy band, transformed into counts per frame, and divided by the pile-up threshold. The thresholds (at which the pile-up level is set to unity) are set to 1.3 counts per frame for MOS and 0.15 counts per frame for pn (Jethwa et al. 2015).

For extended sources, the pile-up level is equal to the measured counts per frame per CCD pixel at the source position divided by the pile-up threshold, and therefore refers to the peak brightness, assuming this can be considered uniform at the pixel scale (4.1" for pn). The threshold is set for all instruments to $5 \times 10^{-3}$ cts per frame per pixel, such that the flux loss is also a few \% when the pile-up level is one.

Among 733 796 point detections, 1171 have PN_PILEUP > 1, among which most (1042) have SUM_FLAG = 1, and only 30 are not flagged (SUM_FLAG = 0). Only 68 detections have PN_PILEUP > 10, among which three are not flagged, all of them in Small Window mode. Similarly, 1388 detections have M1_PILEUP > 1 (22 not flagged) and 1458 have M2_PILEUP > 1 (25 not flagged). All the 167 detections with PILEUP > 10 in any MOS are flagged. The large pile-up values are of course strongly correlated between instruments, and when both are in Full Window mode, MOS is slightly more piled-up than pn (the median ratio of MOS to PN_PILEUP is 1.27). Overall, the number of point-like detections with PILEUP > 1 in any instrument is 2042 (50 not flagged).

6.4.2. Extent likelihood

All detections are tested for their potential spatial extent during the fitting process. The instrumental point-spread function (PSF) is convolved with a beta extent model, fitted to the detection, and the extent likelihood EP_EXTENT_ML is calculated as described by Sect. 4.4.4 of Watson et al. (2009). A source is classified as extended if its core radius (of the beta-model of the PSF), $r_c > 6''$ and if the extended model improved the likelihood with respect to the point source fit such that it exceeded a threshold of $L_{ext, min} = 4$. In the 4XMM catalogues, EP_EXTENT_ML is included for all detections, while it was set to undefined for point-like detections in previous catalogues. $L_{ext, min} \geq 4$ indicates that a source is probably extended, whilst negative values indicate a clear preference of the point-like over the extended fit. As in the previous catalogue, a minimum likelihood difference of four has been chosen to mark a detection as extended. This threshold makes sure that the improvement of the extended over the point-like fit is not only due to statistical fluctuations but from a more precise description of the source profile.
7. The stacked catalogue

A second independent catalogue is compiled in parallel by the XMM-Newton SSC, called 4XMM-DR9s, where the letter "s" stands for stacked. This catalogue lists source detection results on overlapping XMM-Newton observations. The construction of the first version of such a catalogue, 3XMM-DR7s, is described in Traulsen et al. (2019). The construction of 4XMM-DR9s essentially follows the ideas and strategies described there with important changes that are described in full detail in the accompanying paper Traulsen et al. (2020). The two main changes concern the choice of input observations and event-based astrometric corrections before source detection. Also it was found necessary to perform some visual screening of the detections, whose results are reported in the source catalogue.

Observations entering 3XMM-DR7s were filtered rather strictly. Only observations with OBS_CLASS \(< 2\), with all three cameras in full-frame mode, and with an overlap area of at least 20% of the usable area were included. All those limitations were relaxed for the construction of 4XMM-DR9s which resulted in a much larger number of observations to be included and potentially much larger stacks (more contributing observations). Before performing simultaneous source detection on the overlapping observations, individual events were shifted in position using the results from the previous catcorr positional rectification of the whole image processed for 4XMM-DR9. This led to a clear improvement of the positional accuracy in stacked source detection.

All sources found by stacked source detection are listed in 4XMM-DR9s, including those from image areas where only one observation contributes. One may expect some differences between these same sources in 4XMM-DR9 and DR9s, because their input events were treated differently. More information is given in Traulsen et al. (2020).

4XMM-DR9s is based on 1329 stacks (or groups) with 6604 contributing observations. Most of the stacks are composed of two observations, the largest has 352. The catalogue contains 288 191 sources, of which 218 283 have several contributing observations. Auxiliary data products comprise X-ray and optical images, and long term X-ray light curves. Thanks to the stacking process, fainter objects can be detected and 4XMM-DR9s contains more sources compared to the same fields in 4XMM-DR9.

8. Catalogue properties

The 4XMM-DR9 catalogue contains 810 795 detections, associated with 550 124 unique sources on the sky, extracted from 11204 public XMM-Newton observations. Figure 7 shows the distribution of the source fluxes in the total EPIC band and in the soft and the hard band. Also shown in the figure is the distribution of EPIC counts.

Amongst the 4XMM-DR9 detections, 121 792 unique sources have multiple detections, the maximum number of repeat detections being 69, see Fig. 8. 4XMM-DR9 contains 76 999 X-ray detections that are identified as extended objects, that is, with a core radius parameter, \(r_c\), as defined in Sect. 4.4.4 of Watson et al. (2009), \(>6''\) and EP_EXTENT_ML \(\geq 4\), with 74 163 of these having \(r_c < 80''\).

8.1. Astrometry

The systematic astrometric uncertainty of the 4XMM DR9 detection catalogue has been estimated empirically using the SDSS DR14 QSO catalogue (Pâris et al. 2018), following similar steps as those detailed in Sect. 3.3. However, here we use all of the detections in 4XMM-DR9 and any value of poscorrok. The sources in the SDSS DR14 QSO catalogue have been filtered (good quality spectra and avoiding mismatches between target- and matched photometry\(^\text{10}\)). The filtered catalogue has then been cross-matched with the SDSS DR9 photometry catalogue

\(^{10}\) See https://www.sdss.org/dr15/spectro/caveats/
with a maximum distance of 5″. We have discarded all QSOs with more than one SDSS DR9 counterpart out to that distance, keeping only pointlike objects (cL = 6). We cross-correlated the 402,291 “clean” quasars with the “slim” catalogue out to a distance of $r = 30″$. For each of the resulting pairs we have estimated the combined positional error as in Sect. 3.3 and discarded all quasars that had more than one counterpart out to $x = r/\sigma = 6$, making 11,640 suitable quasars (there were 43 quasars with more than one counterpart out to that limit).

Filtering as described in Sect. 3.3 leaves 15,001 quasar-X-ray pairs with $x < 5$. To follow the Rayleigh distribution $P(x) = xe^{-x^2/2}$, we have added an additional positional uncertainty $\Sigma$ in quadrature, so that the total positional uncertainty is now $\sigma' = \sqrt{\sigma^2 + \Sigma^2}$, looking for the value of $\Sigma$ that minimises the log-likelihood of the $x' \equiv r/\sigma'$ and the Rayleigh distribution. We find $\Sigma = 0.961 \pm 0.008$ arcsec for the uncorrected 4XMM-DR9 X-ray positions, where the uncertainty ($1\sigma$) has been calculated by bootstrap with replacement. This can be seen in Fig. 9.

To directly compare the quality of the astrometry in 3XMM-DR8 and 4XMM-DR9, we matched each catalogue of detections with the DR14 release of the SDSS quasar catalogue. Cross-matching was performed without restrictions on the types of XMM-Newton and SDSS sources considered, but we kept only those matches within a matching radius of 15″. This yielded a total of 16,530 3XMM-QSO pairs and 18,002 4XMM-QSO pairs. Figure 10 shows a scatter plot and associated histograms of the RA and Dec offsets between the XMM sources and SDSS quasars. We see that the general astrometric quality of the 4XMM-DR9 catalogue is very good, with mean RA and Dec offsets of $-0.01″$ and $0.005″$ respectively with corresponding standard deviation of $0.70″$ and $0.64″$. No significant improvement is observed when comparing with the 3XMM-DR8–SDSS match.

8.2. Extended sources

Only 76,999 4XMM-DR9 detections (9.5%) are identified as extended, compared to 91,111 in 3XMM-DR8 (11.75% of the catalogue). However, of the extended sources in 4XMM-DR9, 30,464 have the best quality flag (SUM_FLAG = 0, 40% of extended sources), whereas only 12,256 of the 3XMM-DR8 extended sources (13%) have this flag. This implies that the detection of extended sources is more reliable in the new version of the catalogue, with fewer spurious extended sources. This is due to the improved background modelling used for 4XMM-DR9.

9. External catalogue cross-correlation

Cross-correlation with archival catalogues is performed by a distinct pipeline module running at the Observatoire Astronomique de Strasbourg and referred to as the Astronomical Catalogue Data Subsystem (ACDS). For each individual EPIC detection the ACDS lists all possible multi-wavelength identifications located within a $3\sigma$ combined XMM and catalogue error radius from the EPIC position. Finding charts and overlays with ROSAT all-sky survey images of the field are also produced. A detailed description of the ACDS is given in Rosen et al. (2016).

We took the opportunity of the reprocessing of the entire XMM-Newton archive to update the list of archival catalogues and image servers entering the cross-correlation process and finding chart generation. In ACDS version 10.0, a total of 222 catalogues are queried, of which 53 are new with respect to ACDS version 9.0. Among the catalogues providing the largest sky coverage are: GALEX GR6+7 (Bianchi et al. 2017), UCAC4, SDSS DR12 (Alam et al. 2015), panStarrs-DR1 (Chambers et al. 2016), IPHAS DR2 (Barentsen et al. 2014), Gaia DR2 (Gaia Collaboration 2018), 2MASS, AllWISE, Akari, NVSS, FIRST and GLEAM (Hurley-Walker et al. 2017). The XMM-OM Serendipitous Source Survey Catalogue XMM-SUSS4.1 (Page et al. 2012), XMM-Newton slew survey Source Catalogue v. 2.0, the 3XMM-DR8 catalogues, Chandra V2.0 catalogue and the second ROSAT all-sky survey are also queried. Apart from the Chandra Catalogue Release 2.0 whose entries are extracted from the CXC server, all other ACDS catalogues are queried using the Vizier catalogue server.
As for previous releases, 4XMM ACDS tentative identifications are not part of the catalogue proper but are distributed to the community by the XSA and through the XCAT-DB (Michel et al. 2015)\textsuperscript{11}. Finding charts are extracted from several imaging surveys with the following decreasing priority order. First the Sloan digital sky survey (Alam et al. 2015) with colour images made from the g, r and i images extracted from the SDSS server. Second the Pan-STARRS-DR1 (Chambers et al. 2016) with colours images based on the z, g and z+g surveys, third, the MAMA/SRC-J and MAMA/POSS-F plate collections and as a last choice the DSS2 photographic plates. For the one colour photographic surveys, we select the blue image at Galactic latitude \( >20^\circ \), while the red images are preferred in the Galactic plane. Apart from the SDSS, all images are extracted in HEALPix format from Hierarchical Progressive Surveys (HiPS) Aladin server (Fernique et al. 2014).

Spectral energy distributions (SEDs) are provided for each of the unresolved (EP\_EXTENT = 0) unique 4XMM sources. For that purpose, we use basically the same tools as those developed in the framework of the ARCHES project (Motch et al. 2017). The ARCHES algorithm (Pineau et al. 2017) cross-matches in a single pass all selected archival catalogues and for each combination of catalogue entries, computes the cross-match probability. Probabilities are computed from the likelihood that sources in the different catalogues have exactly the same position on the sky, considering their astrometric uncertainties. In particular, the resemblance of the derived SED with that of any given class of objects does not enter in the computation of the probability. The association probability eventually rests on the prior probability that a given X-ray source has a true counterpart in the longer wavelength catalogue considered. This prior is estimated from the observed distribution of X-ray – longer wavelength catalogue associations taking into account the expected rate of spurious matches. In the original ARCHES project, X-ray sources were grouped by XMM observations with similar exposure times, corresponding to similar limiting sensitivities. Although this grouping method offers a clean and relatively easy way to build X-ray source instamles, it still has the disadvantage of mixing bright and faint X-ray sources that will not have the same a priori probability to have a counterpart in the longer wavelength catalogues considered. In order to cope with this potential statistical bias, we designed a method aimed at grouping X-ray sources by range of X-ray flux instead. Accordingly, the ARCHES cross-matching tool had to be modified so as to read the sky area covered by the sample as an input instead of computing it from the list of observations given in entry.

Source detection area requires building EPIC sensitivity maps for each of the XMM observations. In order to compute sensitivity maps, we first tried the approach proposed by for example, Carrera et al. (2007). The method consists of equating the probability of existence of a given source as provided by EP\_8\_DET\_ML with that derived from an excess of counts above a given background assuming Poisson statistics. Although good fits can be obtained for EP\_8\_DET\_ML higher than \(-15\), we found that best fit background areas are highly dependent on off-axis angle and background values when approaching the threshold of EP\_8\_DET\_ML = 6, used as a criteria for a detection to be included in the 4XMM catalogue. Such a discrepancy is not unexpected since the existence probabilities given by the emldetect algorithm also depend on the resemblance of the distribution of photons to that of the PSF. In addition, emldetect relies on the Cash statistics (Cash 1979) and on the approximation of the Wilks theorem to derive probabilities. Instead, we built sensitivity maps by computing at each pixel location the total EPIC broad band count rate that would yield a mathematical expectation of EP\_8\_DET\_ML equal to 6. For that purpose we assume a power law input source spectrum (\( \Gamma = 1.42; N_{\text{H}} = 1.7 \times 10^{20} \text{ cm}^{-2} \)) similar to that of the unresolved sources contributing to the extragalactic background (Lumb et al. 2002). The source spectrum is then folded through the exposure maps and filter responses so as to obtain the source counts in each band and camera in operation. EP\_8\_DET\_ML is then computed taking into account the background maps and the varying shape of the PSF with telescope and off-axis angle.

We estimated the overlap of the 4XMM-DR9 catalogue with 26 archival catalogues selected to cover the largest sky coverage and widest span in wavelengths from UV to radio. The Multi-Order-Coverage map (MOC, Fernique et al. 2015) of each XMM observation was computed with a resolution of 12.8\( \arcmin \) (order 14) and compared to the MOC footprint of each catalogue using a python code developed at CDS (Baumann et al. 2019).

Table 4 lists the pre-selected catalogues sorted by 4XMM coverage. In the optical band, catalogues were prioritised according to their depth, astrometric quality, and range of colours in the following order, SDSS12, PanStars DR1 and Skymapper, so as to cover the entire sky. Whenever a Gaia DR2 match was found within 1.4\( \arcmin \) from the catalogue entry, the Gaia position was assigned to the merged source. APASS9 photometry was added to the merged source if found within a 1.4\( \arcmin \) distance as to extend the photometric measurements to brighter objects. The 1.4\( \arcmin \) radius was derived from the shape of the Rayleigh distribution of the distances between matching sources and guarantees a low rate of false cross-identification. In a similar manner, we cross-matched the ALLWISE and 2MASS catalogues keeping the 2MASS position whenever the difference of position was lower than 3.5\( \arcmin \) at \( |b| \geq 20 \text{deg} \) and 1.5\( \arcmin \) at \( |b| \leq 20 \text{deg} \). Special sky regions such as M31 and the LMC were discarded due to their high optical source density. For each unique source, we only kept the observation offering the highest detection area. 4XMM sources were then grouped into four EPIC (0.2–12.0 keV) ranges of flux with boundaries at 1.4, 3.1 and 7.2 \times 10^{-14} \text{ erg cm}^{-2} \text{ s}^{-1}. This grouping yields a nearly even number of sources in each flux band.

The statistical ARCHES cross-match procedure was applied to five catalogues or group of catalogues: XMM, Galex, SUSS-OM, merged optical and merged infrared. Due to the different areas of the non all-sky catalogues (Galex, SDSS12, PanStars and Skymapper) we split the XMM observations into groups having homogeneous catalogue coverages. In addition, the galactic plane region was treated separately. Finally, a simple cross-match between the ARCHES result and both the AKARI and merged FIRST/NVSS compiled by Mingo et al. (2016) was made. However, their matching likelihoods do not enter in the computation of the overall SED probability provided by the ARCHES tool.

A standard table at CDS\textsuperscript{12} allows us to convert magnitudes into flux. The resulting SEDs are available as individual FITS files and graphical output for the three highest probability SEDs. The sensitivity maps, individual observation MOCS and total 4XMM MOCS are available on the XMM-SSC website\textsuperscript{13}.

\begin{itemize}
\item \textsuperscript{11} \url{http://xcatdb.unistra.fr}
\item \textsuperscript{12} \url{http://vizier.u-strasbg.fr/viz-bin/VizieR-3?-source=METAfltr}
\item \textsuperscript{13} \url{http://xmmssc.irap.omp.eu/Catalogue/4XMM-DR9/4XMM-DR9.html}
\end{itemize}
Table 4. Overlapping area between photometric catalogues and 4XMM observations.

| Catalogue          | Total area covered (deg²) | Overlap with 4XMM (deg²) | Band | Xmatch mode |
|--------------------|---------------------------|--------------------------|------|-------------|
| AllWISE            | all-sky                   | 1152                     | ir   | s           |
| Gaia DR2           | all-sky                   | 1152                     | opt  | s           |
| UCAC4              | all-sky                   | 1152                     | opt  | s           |
| 2MASS              | all-sky                   | 1152                     | ir   | s           |
| APASS              | all-sky                   | 1126                     | opt  | s           |
| Akari              | 39 406                    | 1108                     | farir|x            |
| GMRT               | 36 996                    | 1000                     | radio|x            |
| NVSS               | 34 069                    | 927                      | radio|x            |
| PanStarrs DR1      | 32 134                    | 881                      | opt  | s           |
| GalexGR67          | 26 249                    | 696                      | uv   | s           |
| GLEAM              | 25 423                    | 657                      | radio|x            |
| SkyMapper          | 19 585                    | 550                      | opt  | s           |
| SDSS12             | 14 520                    | 504                      | opt  | s           |
| FIRST              | 10 847                    | 427                      | radio|x            |
| VHS                | 13 670                    | 364                      | ir   | s           |
| XMM-OM-SUSS41      | 348                       | 343                      | uv   | s           |
| SUMSS              | 8354                      | 216                      | radio|x            |
| UKIDSS LAS         | 3695                      | 174                      | ir   | s           |
| VST                | 3988                      | 86                       | opt  | s           |
| Galex MIS          | 1880                      | 83                       | uv   | s           |
| VPHAS              | 670                       | 77                       | opt  | s           |
| UKIDSS GPS         | 1366                      | 76                       | ir   | s           |
| WBH2005 20         | 614                       | 72                       | radio|x            |
| Glimpse            | 471                       | 70                       | ir   | s           |
| IPHAS              | 1888                      | 59                       | opt  | s           |
| WBH2005 6          | 164                       | 35                       | radio|x            |

Notes. The last column shows the way the catalogue was processed, either using the ARCHES multi catalogue statistical cross-match (s) or using a simple positional cross-match (x).

10. Catalogue access

The catalogue of detections is provided in several formats. A Flexible Image Transport System (FITS) file and a comma-separated values (CSV) file are provided containing all of the detections in the catalogue. For 4XMM-DR9 there are 810 795 rows and 336 columns. A separate version of the catalogue (the slim catalogue) with only the unique sources is provided, that is, 550 124 rows, and has 45 columns, essentially those containing information about the unique sources. This catalogue is also provided in FITS and CSV format. We also provide SQL CREATE statements to load the data in CSV format. These can be found on the XMM-Newton Survey Science Centre webpages16. The stacked catalogue is provided in FITS format only. Ancillary tables to the catalogue also available from the XMM-Newton Survey Science Centre webpages include the table of observations corresponding to the Xmatch catalogue. The XMM-Newton Survey Science Centre webpages provide access to the 4XMM catalogue, as well as links to the different servers distributing the full range of catalogue products. These include, the ESA XMM-Newton archive (XSA), which provides access to all of the 4XMM data products, and the ODF data, the XCat-DB15 produced and maintained by the XMM-Newton SSC, which contains possible EPIC source identification produced by the pipeline by querying 222 archival catalogues, see Sect. 9. Finding charts are also provided for these possible identifications. Other source properties as well as images, time series and spectra are also provided. Multi-wavelength data taken as a part of the XID (X-ray identification project) run by the SSC over the first fifteen years of the mission are also provided in the XIDresult database16. The XMM-SSC catalogue server17 provides access to each source and regroups information concerning all of the detections for a unique source. It also provides the XMM-Newton lightcurves and spectra and permits the user to undertake simple spectral fitting, as well as overlays of the same region of sky in all wavelengths. The catalogue can also be accessed through HEASARC18 and VIZIER19. The results of the external catalogue cross-correlation carried out for the 4XMM catalogue (Sect. 9) are available as data products within the XSA or through the XCat-DB, The XMM-Newton Survey Science Centre webpages also detail how to provide feedback on the catalogue.

Where the 4XMM catalogue is used for research and publications, please acknowledge their use by citing this paper and including the following:

“This research has made use of data obtained from the 4XMM XMM-Newton serendipitous source catalogue compiled by the ten institutes of the XMM-Newton Survey Science Centre selected by ESA.”

We note that the 4XMM catalogue of detections, as for previous versions of this catalogue, contains detections with a significance as low as ~3σ (Maximum likelihood of six), along with sources that have been flagged as possibly spurious. Statistically some of these sources will be spurious. In order to create the cleanest catalogue possible, where statistically almost all sources are real, it is necessary to filter the catalogue to include only EPIC sources with for example, a 5σ significance (Maximum likelihood of ~14) and to keep only those with with no flags, for example, EP_8_DET_ML > 14 && SUM_FLAG < 1

Filtering with these criteria for 4XMM-DR9 leaves 433 612 detections. 99.6% or 431 924 of the point-like detections have no pileup (XX_PILEUP < 1, where XX is either pn, M1 or M2 for the pn, MOS 1 or the MOS 2 detectors).

11. Upper limits for observed regions of the sky

The XMM-SSC provides an upper limit server for the user to determine an upper limit for the flux given a non-detection in a region observed by XMM-Newton. The server is known as FLIX (Flux Limits from Images from XMM-Newton). This upper limit can be calculated for any of the standard XMM-Newton bands for a user defined statistical significance and sky region. A single region or many regions may be queried at the same time. This upper-limit flux is determined empirically using the algorithm described by Carrera et al. (2007). A link to the FLIX upper limit server is provided on the XMM-SSC webpages and the ESA SOC webpages20.
12. Limitations of the catalogue

In order for the data processing to work successfully for the 14,041 observations that were reduced in the production of the catalogue, some limitations to the full parameter space available for some values are imposed. These are described in the following sections.

12.1. Maximum extent of extended detections

When dealing with extended detections, the software determines the radius of the detection, up to a limit of 80″ to optimise processing time. Whilst this may appear restrictive, only 0.007% of the catalogue detections are clean and extended, with a radius of >80″.

12.2. Error values on counts, rate and flux

Should a detection fall close to a chip gap or the edge of the field of view on one or more cameras, only a small fraction of the point spread function will be recorded for that camera. The fraction is given by the XX_MASKFRAC columns, where XX refers to EP (EPIC), PN (pn), M1 (MOS 1) or M2 (MOS 2). Where the XX_MASKFRAC value is low, the error on the counts, rate or flux may be very high, compared to the value of the counts, rate or flux, as these quantities are derived for the whole PSF. We note that detections which have less than 0.15 of their PSF covered by the detector are considered as being not detected.

13. Future catalogue updates

Incremental releases (data releases) are planned to augment the 4XMM catalogue. At least one year of data will be included with each data release. Data release ten (DR10) will provide data becoming public during 2019 and should be released during 2020. These catalogues will be accessible as described in Sect. 10.

14. Summary

This paper describes the improvements made to the software and calibration used to produce the new major version of the XMM-Newton catalogue, 4XMM. The 4XMM-DR9 catalogue contains 810,795 detections in the X-ray band between 0.2 and 12.0 keV. The catalogue covers 1152 degrees^2 (2.85%) of the sky. In terms of unique X-ray sources, the 4XMM-DR9 catalogue is the largest X-ray catalogue produced from a single X-ray observatory, with 550,124 unique sources compared to 317,167 unique X-ray sources in the Chandra source catalogue v. 2.0 and 206,335 unique X-ray sources in the 2XPS catalogue of X-ray sources from the Neil Gehrels Swift Observatory. In this new version of the catalogue, source detection has been shown to be much improved, with fewer spurious sources and in particular, many fewer spurious extended sources. In addition, we provide lightcurves and spectra for a much larger fraction of the catalogue than in previous versions (36% of detections in 4XMM-DR9 compared to 22% of detections in 3XMM-DR8). These spectra and lightcurves benefit from finer binning (MOS spectra and pn lightcurves). The catalogue benefits from extra complementary products, such as multi-wavelength spectral energy distributions for each source, sensitivity maps and catalogue footprint maps. We provide information on how to access the catalogue as well as how to retrieve upper limits for non-detections in the catalogue footprint. The catalogue is ideal for quick access to data products (fluxes, spectra, images, etc), searching for new objects, population studies of homogeneous samples and cross correlation for multi-wavelength studies.

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