Trigonometric Parallaxes of 6.7 GHz Methanol Masers

K. L. J. Rygl1,⋆, A. Brunthaler1, M. J. Reid2, K. M. Menten1, H. J. van Langevelde3,4 and Y. Xu5

1 Max-Planck-Institut für Radioastronomie (MPIfR), Auf dem Hügel 69, 53121 Bonn, Germany
e-mail: [kazi;brunthal;kmenten]@mpifr-bonn.mpg.de
2 Harvard-Smithsonian Center for Astrophysics, 60 Garden Street, Cambridge, MA 02138, USA
e-mail: reid@cfa.harvard.edu
3 Joint Institute for VLBI in Europe, Postbus 2, 7990 AA Dwingeloo, the Netherlands
e-mail: langevelde@jive.nl
4 Sterrewacht Leiden, Leiden University, Postbus 9513, 2300 RA Leiden, the Netherlands
5 Purple Mountain Observatory, Chinese Academy of Sciences, Nanjing 21008, China
e-mail: xuye@pmo.ac.cn

Preprint online version: October 2, 2009

ABSTRACT

Aims. Emission from the 6.7 GHz methanol maser transition is very strong, relatively stable, has small internal motions and is observed toward numerous massive star forming regions in the Galaxy. Our goal is to perform high precision astrometry using this maser transition to obtain accurate distances to their host regions.

Methods. Eight strong masers were observed during five epochs of VLBI observations with the European VLBI Network between 2006, June, and 2008, March.

Results. We report trigonometric parallaxes for five star forming regions, with accuracies as good as ~ 22 µas. Distances to these sources are 2.57±0.34 kpc for ON 1, 0.776±0.034 kpc for L 1206, 0.929±0.034 kpc for L 1287, 2.38±0.13 kpc for NGC 281-W and 1.59±0.06 kpc for S 255. The distances and proper motions yield the full space motions of the star forming regions hosting the masers, and we find that these regions lag circular rotation on average by ~17 km s⁻¹, a value comparable to that found recently by similar studies.

Key words. Techniques: interferometric – Astrometry – Masers – Stars: formation – ISM: molecules – Galaxy: kinematics and dynamics

1. Introduction

Accurate distances and proper motions are crucial for studies of the structure and kinematics of the Milky Way. Since we are in the Galactic plane, it is not easy to determine the spiral structure of our Milky Way (Reid et al. 2009c). Massive star forming regions trace the spiral arms and are well suited objects to reveal the structure of the Milky Way.

The determination of fundamental physical properties of individual objects, such as size-scales, masses, luminosities and ages also depend critically on distance. For example, the distance to the Orion Nebula, determined by a number of trigonometric parallax measurements, using radio continuum emission from stars in its associated cluster (Menten et al. 2007; Sandstrom et al. 2007), water masers (Hirota et al. 2007) and SiO masers (Kim et al. 2008), turned out to be 10% smaller than previously assumed, resulting in 10% lower masses, 20% fainter luminosities and 20–30% younger ages for the stars in the cluster.

The most fundamental and unbiased method to measure distances is the trigonometric parallax, which depends only on geometry and is therefore free of any astrophysical assumptions. To achieve the sub-milliarcsecond astrometric accuracy at optical wavelengths requires space-borne observations. The first dedicated optical satellite for this purpose was ESA’s Hipparcos (Perryman et al. 1995). Hipparcos, with its accuracy of 0.8–2 mas, was capable of measuring distances up to ~200 pc – only a small fraction of the Milky Way. A new optical astrometry satellite GAIA, to be launched in 2012, will be two orders of magnitude more accurate, and push the optical parallaxes to Galactic size-scales (e.g., Lindegren 2009). Despite the high accuracy, GAIA will suffer from dust extinction in the Galactic plane, in particular in the spiral arms and toward the Galactic center. Here radio astronomy can provide an important complement to GAIA; radio wavelengths do not suffer from dust extinction and Very-Long-Baseline Interferometry (VLBI) phase-referencing techniques can provide accuracies better than 10 µas, allowing accurate distances with errors less than 10% out to 10 kpc (see for example Honma et al. 2007; Hachisuka et al. 2009; Reid et al. 2009c).

Strong and compact radio sources, such as molecular masers, are ideal targets for trigonometric parallax measurements. Masers are frequently found in (dusty) star formation regions (SFRs) and asymptotic giant branch stars. The most numerous interstellar masers are the 22 GHz water masers, followed in number by the 6.7 GHz methanol masers; to date several hundreds of these masers have been found across the Galaxy (e.g., Pestalozzi et al. 2005; Green et al. 2009). They are exclusively associated with early stages of massive star formation, observed both prior to the development of an ultra compact H ii (UCH ii) region and coexistent with these (Menten 1991).
Table 1. Observation information

| Program | Date       | Participating Antennas |
|---------|------------|------------------------|
| EB32A   | 10 Jun 2006| Jb Wb Ef Mc Tr Nt Hh On Da |
| EB32B   | 18 Mar 2007| Jb Wb Ef Mc Tr Nt Hh On Cm |
| EB32C   | 17 Jun 2007| Jb Wb Ef Mc Tr Nt Hh On Cm |
| GB63A   | 31 Oct 2008| Jb Wb Ef Mc Tr Hh On Y1  |
| GB63B   | 16 Mar 2008| Jb Wb Ef Mc Tr Nt Hh On Y1 |

Notes: a) The full names of the antennas are: Jb – Jodrell Bank; Wb – Westerbork (in single dish mode); Ef – Effelsberg; Mc – Medicina; Tr – Torun; Nt – Noto; Hh – Hartebeesthoek; On – Onsala; Da – Darnhall; Cm – Cambridge; Y1 – single EVLA antenna.

Ellingsen (2006), Pestalozzi et al. (2007). Maser modeling indicates that the 6.7 GHz emission is likely excited via radiative pumping by warm dust heated by newly formed high-mass stars (Sobolev et al. 2007).

Sources of methanol masers have long lifetimes of ~ 10^4 year (van der Walt 2005). Their velocity spread is typically within ±5 km s^{-1} about the systematic velocity of the molecular core, and single maser features are usually narrow (~1 km s^{-1}). The kinematics suggest that methanol masers originate from different regions than water masers, which have wider velocity spreads and are produced in proto-stellar outflows (Menten 1996). To summarize, the strength, ubiquity, long lifetimes and small internal motions make the 6.7 GHz methanol maser very suitable for astrometric purposes.

Here we present results from observations with the European VLBI Network (EVN) of eight strong methanol masers belonging to well-known massive SFRs in the outer part of the Milky Way: ON 1, L 1287, L 1206, NGC 281-W, MonR 2, S 252, S 255 and S 269. Recently, H_2O maser or 12.2 GHz methanol maser parallax measurements have been reported for three of these regions (NGC 281-W, S 252, Reid et al. (2009a); S 269, Honma et al. (2007)), allowing a valuable cross check with our measurements. The results of this present work are crucial for star formation studies in these regions, and, together with a larger sample of parallaxes, will help to understand the structure of the Local and Perseus arms.

In this paper we report on the first parallax measurements made with the EVN. Results presented here replace preliminary, less accurate and less complete measurements (Rygl et al. 2008). In the next section we describe the observations and data analysis and in Sect. 3 we explain the method and fitting of the parallax and proper motion. The results are then presented per source in Sect. 4 and discussed in Sect. 5. The conclusions are summarized in Sect. 6.

2. Observations and data analysis

The observations were performed with the EVN at five epochs between 2006, June, and 2008, March. The exact dates of the observations and the participating antennas are listed in Table 1. Each observation lasted 24 hours and made use of geodetic-like observations to calibrate the tropospheric zenith delays at each antenna (see Reid & Brunthaler 2004; Brunthaler et al. 2005; Reid et al. 2009a, for a detailed discussion). A typical observing run consisted of four 6-hour blocks containing ~1 hour of geodetic-like observations, ~10 minutes of observation of fringe finders. The remaining time was spent on maser/background source-phase-referencing observations. During each run, the average on-source time per maser was between ~0.9 and ~1.2 hours depending on the sky position; half of our targets were low declination sources and had a limited visibility.

Using the technique of phase-referencing, each maser was observed in a cycle with two (or one) nearby (~1°–2° separation) compact radio quasars, which were used as background sources. The sources were switched every 2 minutes. Before our EVN observations, we had used the NRAO Very Large Array under project AB1207 in A-configuration to observe several compact NVSS (Condon et al. 1998) sources within 1° from the maser source at two frequencies (8.4 and 4.8 GHz) to check their spectral index and compactness. For the best candidates, we determined their position to a sub-arcsecond accuracy to use them as a position reference in the EVN observations. Additionally, several known VLBA calibrators were used, namely: J2003+3034 and J0613+1708 (Fomalont et al. 2003), J0603+2155(S9) and J0613+1306 (Ma et al. 1998), J0047+5657 (Beasley et al. 2002), J2223+6249 (Petrov et al. 2005), J0035+6130 (Petrov et al. 2006) and J0606-0724 (Kovalev et al. 2007). Table 2 lists the positions of the masers and their background sources.

The observations were performed with eight intermediate frequency bands (IF) of 8 MHz width each in dual circular polarization sampled at the Nyquist rate with 2 bits per sample, yielding a recording rate of 512 Mbps. The data were correlated in two passes at the Joint Institute for VLBI in Europe (JIVE), using an integration time of 0.5 seconds, affording a field of view of 1.2 (limited by time averaging smearing). The maser data were correlated using one 8 MHz IF band with 1024 spectral channels, resulting in a channel separation of 7.81 kHz or 0.41 km s^{-1} at 6.7 GHz. The quasar sources were correlated in continuum mode with eight IFs of 8 MHz width with a channel separation of 0.25 MHz.

The data reduction was done using the NRAO Astronomical Image Processing System (AIPS). The geodetic-like observations were reduced separately, and tropospheric delays were estimated for each antenna. The data were reduced following the EVN guidelines, applying parallactic angle and ionospheric delay corrections. The JIVE correlator model uses Earth’s orientation parameters which are interpolated from the appropriate daily-tabulated values, hence it is not necessary to correct them after the correlation. The ionospheric delays were based on the JPL GPS–IONEX total vertical electron content maps of the atmosphere. Amplitudes were calibrated using system temperature measurements and standard gain curves. A “manual phase-calibration” was performed to remove delay and phase differences between the IFs. The correction for the Earth rotation was done with the task ‘CVEL’. For each maser, a spectral channel with one bright and compact maser spot was used as the phase reference. The data was Hanning-smoothed to minimize Gibbs ringing in the spectral line data. To avoid the strong fluctuations caused by the bandpass edges, the outer two channels in every IF were discarded (following Reid et al. 2009a). The positions of the masers and background sources were extracted by fitting 2D Gaussians to the maps.

Here, we report the results on the first five sources for which we have completed the analysis: ON 1, L 1206, L 1287, NGC 281-W and S 255. For the other sources, S 252, S 269 and MonR 2, we had problems with the calibration of the data which were likely caused by residual atmospheric delay, and hence we were unable to achieve sufficient accuracy for a parallax measurement.
As a result, di-

ference and phase shifts from the old EVN apparent position to


time-consuming with available software because

fringe rate which could not be interpolated correctly when ap-

portions, comparable to the residual tropospheric delay at the

observing frequency. Our single frequency geodetic-like obser-

ments of the total electron content of the ionosphere obtained

Therefore we choose to discard some of the first-epoch data.

Most of the ionospheric delay is removed using measure-

ments of the total electron content of the ionosphere obtained

from dual-band global positioning system measurements (e.g.,

Ros et al. 2000). The expected residual delay is still a few cen-

imeters, comparable to the residual tropospheric delay at the

observing frequency. Our single frequency geodetic-like observ-

ations do not allow a separation of the ionospheric and tro-

pospheric contributions. The measured group-delays from the

gedeotic blocks are interpreted as of tropospheric origin, and

the data are phase-delay-corrected accordingly. An ionospheric

group-delay contributes differently to the phase-delay, and there-

fore the tropospheric correction can even deteriorate the calibra-

tion when the group delay has partially an ionospheric origin.

Since the residual tropospheric delay is a few centimeters at 6

GHz and it can easily be confused with an ionospheric resid-

ual, the tropospheric correction is not expected to be very useful

in improving the quality of calibration. As a test, all data were

duced with and without a “tropospheric” delay correction. In

ly more than half of the cases the signal-to-noise ratio im-

proved by applying the tropospheric correction. We noted a trend

of improvement with decreasing declination: both high declina-

sources, δ ≃ 64°. L 1206 and L 1287 improved in 2 out of 5

epochs, NGC 281-W at δ ≃ 57° improved in 3 out 5 epochs and

ON 1 at δ ≃ 31° even improved in 4 out of 5 epochs. Ionospheric

delay saturates at zenith angles greater than ≃ 70°. However, the
tropospheric delay continues to grow rapidly at larger zenith an-
gles and can dominate over the ionospheric delay at zenith an-
gles of ≃ 80° (see Thompson et al. 1991).

For each SFR, we found the emission to arise from a num-

ber of separate maser spots. Most of the line profiles stretched

over several channels. Both spatially and/or spectrally different

maser spots were considered as distinct maser components. We

pected the behavior of the proper motion for each maser spot

relative to the reference spot. Maser spots with strong non-linear

proper motions or a large scatter of position about a linear fit

were discarded. Only compact maser spots with well behaved

proper motions were used for the parallax fitting.

The average internal proper motion of the maser spots ranged

between 0.06 and 0.23 mas yr−1. Considering the distance of

each maser, these proper motions correspond to 0.5–1 km s−1,
much lower than the internal proper motions of water masers,
which can reach up to 20–200 km s−1 (Hachisuka et al. 2006).
The only exception was ON 1, which separates into two distinct

| Source | R.A. (J2000) (h:m:s) | Dec. (J2000) (° : ’ :") | φ° (∆) | P.A.° (∆) | Brightness° (Jy beam−1) | Restoring Beam° (mas, mas, deg) |
|--------|---------------------|--------------------------|--------|-----------|------------------------|-------------------------------|
| ON 1    | 20:10:09.074        | +31:31:35.946            | ...    | 1–7       | 5.7 × 3.9, 86          |
| J2003+3034 | 20:03:30.244        | +30:34:30.789            | 1.71   | -124      | 0.124                  |
| J2009+3049 | 20:09:17.588        | +30:49:24.580            | 0.73   | -14       | 0.008                  |
| L 1206  | 22:28:51.407        | +64:13:41.314            | ...    | 0.3–1.4   | 5.7 × 4.0, 80          |
| J2223+6249 | 22:23:18.097        | +62:49:33.805            | 1.53   | -157      | 0.093                  |
| J2225+6411 | 22:25:27.993        | +64:11:15.030            | 0.37   | -96       | 0.002                  |
| L 1287  | 00:36:47.353        | +63:29:02.162            | ...    | 0.5–4.1   | 6.0 × 4.1, 57          |
| J0035+6130 | 00:35:25.311        | +61:30:30.761            | 1.98   | -176      | 0.092                  |
| J0037+6236 | 00:37:04.332        | +62:36:33.310            | 0.88   | 178       | 0.061                  |
| NGC 281-W | 00:52:24.196        | +56:33:43.175            | ...    | 1–12      | 5.7 × 4.2, 56          |
| J0047+5657 | 00:47:00.429        | +56:57:42.395            | 0.84   | 62        | 0.142                  |
| J0052+5703 | 00:52:54.303        | +57:03:31.460            | 0.50   | -8        | 0.010                  |
| MonR2   | 06:07:47.862        | -06:22:56.518            | ...    | 9–22      | 8.6 × 4.4, 42          |
| J0606-0724 | 06:06:43.546        | -07:24:30.232            | 1.06   | -165      | 0.135                  |
| S 252   | 06:08:53.344        | +21:38:29.158            | ...    | 17–44     | 6.4 × 3.9, 56          |
| J0603+2158 | 06:03:51.557        | +21:59:37.698            | 1.22   | -73       | 0.019                  |
| S 255   | 06:12:54.020        | +17:59:23.316            | ...    | 4         | 6.9 × 3.9, 56          |
| J0613+1708 | 06:13:36.300        | +17:08:24.946            | 0.87   | 169       | 0.040                  |
| S 269   | 06:14:37.055        | +13:49:36.156            | ...    | 0.2–1.1   | 6.8 × 3.9, 58          |
| J0613+1306 | 06:13:57.693        | +13:06:45.401            | 0.73   | -167      | 0.087                  |

a Separation, φ, and the Position Angle (East of North), P.A., between the maser and the background source.
b The brightness and restoring beam (East of North) are listed for the second epoch.
maser groups with a relative proper motion of 0.52 mas yr\(^{-1}\), or 6.3 km s\(^{-1}\). This particularity is discussed in Sect. 3.1.

The parallaxes and proper motions were determined from the change in the positions of the maser spot(s) relative to the background source(s). The data were fitted with a parallax and a linear proper motion. Since the formal position errors are only based on the signal-to-noise ratios determined from the images, they do not include possible systematic errors from residual zenith delay errors or source structure changes. This leads to a large reduced \(\chi^2\) value of the fits. Hence, we added additional error floors in quadrature to the positions until reduced \(\chi^2\) values close to unity were reached for each coordinate.

First, we performed parallax and proper motion fits for each maser spot relative to one background source. Then we made a combined fit, in which all the maser spots relative to one background source were used, assuming one parallax but different proper motions for each maser spot. Finally, we repeated this combined fit for both the background sources together. The position measurements of different maser spots are not independent, since systematic errors, such as an unmodeled atmospheric delay, will affect all maser spots in a similar way. If these systematic errors dominate, this will lead to unrealistically small errors. The most conservative approach, which we adopted, is to assume that the systematic errors are 100% correlated. Then the error of the combined fit has to be multiplied by \(\sqrt{N}\), where \(N\) is the number of maser spots.

However, this will overestimate the error, if significant random errors are present (e.g., owing to maser blending and structural changes over time), since the latter are not correlated between different maser spots. Random errors can be reduced by averaging the positions of the different maser spots (following the approach of Bartkiewicz et al. 2008, Hachisuka et al. 2009).

We calculated the average positions with respect to each background source after removing their position offsets and proper motions. Then, we performed a parallax fit on these averaged data sets relative to each individual background source, and both the background sources combined. This approach has the advantage that we can reduce the random errors, while leaving the systematic errors intact. In Table 3, we list, for each maser, the individual parallax and proper motion fits to each maser spot, the combined fit of all the maser spots and the fit of the averaged data sets.

For most masers we have observed two background sources. Some of these background source pairs show a variation in their separation (mimicking a proper motion) of up to 1–2 mas yr\(^{-1}\) (on average this apparent movement was 0.54 mas yr\(^{-1}\)). This apparent movement is much larger than expected for extragalactic sources measured at 12 GHz, < 0.1 mas yr\(^{-1}\) (Reid et al. 2009) or 22 GHz, < 0.02 mas yr\(^{-1}\) (Brunthaler et al. 2007). While at higher frequencies, the radio emission is typically dominated by the flat spectrum radio cores, the steep spectrum emission from the radio jets becomes more important at lower frequencies. Thus, structure changes in these jets can lead to apparent motions, even if these jets are unresolved. At the distances of the respective masers, these apparent movements correspond to 2–10 km s\(^{-1}\). Since we do not know which of the two background sources is responsible for this movement, the apparent movements of the background source pairs were added in quadrature to the final errors of the averaged proper motions of the masers.

A summary of our parallaxes and proper motions is given in Table 3.

### Table 3. Parallax and proper motion results

| Source | \(\pi\) (mas) | \(D_e\) (kpc) | \(\mu^\pi_a\) (mas yr\(^{-1}\)) | \(\mu^\pi_b\) (mas yr\(^{-1}\)) |
|--------|-------------|-------------|-----------------------------|-----------------------------|
| ON 1   | 0.389 ± 0.045 | 2.57\(^{+0.34}_{-0.27}\) | -3.24 ± 0.89\(^{-9}\) | -5.42 ± 0.46\(^{-9}\) |
| L 1206 | 1.289 ± 0.153 | 0.776\(^{+0.104}_{-0.080}\) | 0.27 ± 0.23 | -1.40 ± 1.95 |
| L 1287 | 1.077 ± 0.039 | 0.929\(^{+0.034}_{-0.033}\) | -0.86 ± 0.11 | -2.29 ± 0.56 |
| NGC 281-W | 0.421 ± 0.022 | 2.38\(^{+0.13}_{-0.12}\) | -2.69 ± 0.16 | -1.77 ± 0.11 |
| S 255  | 0.628 ± 0.027 | 1.59\(^{+0.07}_{-0.06}\) | -0.14 ± 0.54 | -0.84 ± 1.76 |

*\(^a\) The proper motion, \(\mu^\pi_a\), includes the factor \(\cos(\delta)\). |
*\(^b\) The errors of the proper motion take the uncertainty resulting from an apparent nonzero movement between the background sources into account. |
*\(^c\) For the proper motion of ON 1, we took an average of the north and south component.

### 4. Individual sources

#### 4.1. Onsala 1

The maser emission of Onsala 1 (ON 1) consists of two groups separated spatially by ~940 mas which have different radial velocities by 15 km s\(^{-1}\) (Fig. 1); the northern group, which includes the reference channel, is centered at \(V_{LSR} = 0\) km s\(^{-1}\) and the southern group at \(\pm 15\) km s\(^{-1}\). Four maser spots in the northern group were suitable for parallax fitting. The masers in the southern group were not used for the parallax fit, since the phase-calibration was less accurate because of the large distance to the phase-reference center (located in the northern maser group). We find a parallax of 0.389 ± 0.045 mas, corresponding to a distance of 2.57\(^{+0.34}_{-0.27}\) kpc. The proper motions of the southern maser spots were fitted by assuming the parallax result of the northern group. All results are listed in Table 3 and the parallax fit is plotted in Fig. 5.

For ON 1, two background sources were detected in epochs two through five. The last (fifth) epoch had bad \((u,v)\) coverage, which resulted in very elongated synthesized beams in the east-west direction (J2003+3034, 12.1 \(\times\) 3.7 mas\(^2\), P.A. ~ 87\(^\circ\); J2009+3049, 11.0 \(\times\) 3.3 mas\(^2\), P.A. ~ 88\(^\circ\)) compared to the representative numbers listed in Table 2. Hence, the right ascension data had a large uncertainty in epoch five. We estimated this uncertainty from fitting the variation in the background-source pair separation to be 2 mas (see Fig. 4). Since this was a combined error for both the background sources, the position error for each individual background source was \(\sqrt{2}\) mas. We added an additional error floor of \(\sqrt{2}\) mas to the error given for all the right ascension data points in the fifth epoch. The apparent movements between the two background sources, J2003+3034–J2009+3049, were 0.88 ± 0.06 mas yr\(^{-1}\) in right ascension and 0.45 ± 0.21 mas yr\(^{-1}\) in declination (Fig. 4).

#### 4.2. L 1206

In L 1206, two maser groups were found separated by ~100 mas. A third maser group was found north-east of the other two. This third maser group was not detected in all the epochs due to its weakness and was not used in the parallax measurements. Also the maser spot at ~11.6 km s\(^{-1}\) was omitted for the same reason. The masers in L 1206 are shown together with a spectrum in Fig. 5. The parallax fit used four compact maser spots, and resulted in 1.289 ± 0.153 mas, corresponding to a distance of...
0.776^{+0.104}_{-0.083}\text{ kpc. The results of the parallax and proper motion fits are displayed in Fig. 7 and listed in Table 5.}

Only one of the two background sources, namely J2225+6411, was detected in the first epoch. The other background source, J2223+6249, had a much larger separation from the maser (i.e., the phase reference) and likely the transfer of the phase solutions to the J223+6249 data failed. The apparent movements between the background sources were 0.22 ± 1.45 mas yr$^{-1}$ in right ascension and 1.91 ± 1.43 mas yr$^{-1}$ in declination (Fig. 8).
Fig. 3. Results of the parallax fit based on four maser spots. The left graph shows a combined fit on all data, the right graph is a fit on the averaged data sets. The filled colored dots mark the data points in right ascension, while the filled colored triangles mark the declination. The solid line is the resultant fit in right ascension, the dashed line in declination. Different colors indicate a different background source.

Fig. 7. Results of the parallax fit for L 1206 based on four maser spots. The left graph shows a combined fit on all data, the right graph is a fit on the averaged data sets. The filled dots mark the data points in right ascension, while the filled triangles mark the declination. The solid line is the resultant fit in right ascension, the dashed line in declination. Different colors indicate different background sources.

4.3. L 1287

Toward L 1287 we find three maser groups, two close together (within \( \sim 40 \) mas) and the third \( \sim 135 \) mas southward, shown together with a spectrum in Fig. 9. A total of six maser spots from the three groups were used for the parallax fit. We report a parallax value of \( 1.077 \pm 0.039 \) mas, corresponding to a distance of \( 0.929^{+0.034}_{-0.033} \) kpc. The results are plotted in Fig. 11 and listed in Table 5.

The variation of the separation between the two background sources, J0035+6130 and J0037+6236, is shown in Fig. 12. Two data points appear to be outliers: the right ascension at epoch
four, and declination at epoch two. Increasing the data errors to 0.9 and 0.7 mas, respectively, for these outliers yields an apparent movement fit of $-0.05 \pm 0.12$ mas yr$^{-1}$ in right ascension and $0.54 \pm 0.10$ mas yr$^{-1}$ in declination (Fig. 13). We investigated the visibility amplitudes as a function of $(\mu, \nu)$ distance and found evidence for extended structure. For J0037+6236 at the fourth epoch, evidence of extended structure in the right ascension direction was found, however, for J0035+6130 at the second epoch, we found no indications of such structure.

### 4.4. NGC 281-W

We found four groups of methanol masers toward NGC 281-W, as illustrated in Fig. 14, together with the source spectrum. For the parallax fit we used a total of six maser spots originating from three groups, the central, southern and south-east group (Fig. 13). The weaker maser spots, like the ones located in the western group were omitted in the parallax fitting. We found four groups of methanol masers toward NGC 281-W, as illustrated in Fig. 13 together with the source spectrum. We investigated the visibility amplitudes as a function of $(\mu, \nu)$ distance and found evidence for extended structure. For J0037+6236 at the fourth epoch, evidence of extended structure in the right ascension direction was found, however, for J0035+6130 at the second epoch, we found no indications of such structure.

### 4.5. S 255

For S 255, we could use only emission from one maser channel, namely the reference channel. Based on this maser spot, at 4.6 km s$^{-1}$ (Fig. 17), we find a parallax of $0.628 \pm 0.027$ mas corresponding to a distance of $1.59^{+0.07}_{-0.06}$ kpc. The results are shown in Fig. 17 and listed in Table 5.

We observed only one background source, J0613+1708, for S 255 (Fig. 18), which was detected in epochs two to five. With only one background source, we could not check that its apparent movement was small, and so, conservatively, we added uncertainties of 1 mas yr$^{-1}$ in both coordinates.

## 5. Discussion

### 5.1. Space motions

With parallax and proper motion measurements, one can calculate the full space motion of the masers in a SFR in the Galaxy. Using an accurate model of Galactic dynamics (Galactic rotation speed, $\Theta_0$, distance of the Sun to the Galactic center, $R_0$, rotation curve) for a source, and removing the modeled contribution of Galactic rotation, allows one to retrieve source peculiar motion relative to a circular orbit. The peculiar motion is described by the vectors $U$, $V$, and $W$, locally toward the Galactic center, toward the direction of rotation and toward the North Galactic Pole, respectively.

In Table 4 the peculiar motions of our sources are given for two different Galactic models: the I.A.U. recommended values for the LSR motion $R_0 = 8.5$ kpc, $\Theta_0 = 220$ km s$^{-1}$ in columns 2–4; a new model with $R_0 = 8.4$ kpc, $\Theta_0 = 254$ km s$^{-1}$, based on a large sample of trigonometric parallax measurements (Reid et al. 2009b), in columns 5–7. Both models adopt the Hipparcos values for the Solar motion (Dehnen & Binney 1998) and assume a flat rotation curve. We used the parallaxes
Fig. 15. Results of the parallax fit for NGC 281-W based on six maser spots. The left graph shows a fit on all data, while the right graph is a fit on the averaged data sets. The filled dots mark the data points in right ascension, while the filled triangles mark the declination. The solid line is the resultant fit in right ascension, the dashed line in declination. Different colors indicate a different background source.

Fig. 16. Left: Variation of the separation between background sources J0047+5657–J0052+5703 belonging to NGC 281-W. Right: the proper motion fit on the component of J0052+5703 with respect to the central source J0052+5703. The solid line and the dots represent the right ascension data, while the dashed line and the filled triangles represent the declination data.

Table 4. Peculiar motions

| Source      | $R_0 = 8.5$ kpc, $\Theta_0 = 220$ km s$^{-1}$ | $R_0 = 8.4$ kpc, $\Theta_0 = 254$ km s$^{-1}$ |
|-------------|---------------------------------------------|---------------------------------------------|
|             | $U$  | $V$  | $W$  | $U$  | $V$  | $W$  | $U$  | $V$  | $W$  |
| ON 1        | 18±3 | −19±3| 4±10 | 7±8  | −21±3| 4±10 |
| L 1206      | 2±4  | −16±2| 0±6  | −1±4 | −16±2| 0±6  |
| L 1287      | 13±2 | −18±2| −3±2 | 10±2 | −18±2| −3±3 |
| NGC 281-W   | 13±2 | −3±2 | −9±2 | 6±2  | −3±2 | −9±2 |
| S 255       | 1±3  | −4±12| 3±7  | 2±3  | −4±12| 3±7  |

and proper motions as are listed in Table 3, where the uncertainty introduced by the apparent movement of the background sources was included in the proper motion uncertainty of the source.

In Fig. 20 we plot the SFRs studied in this work at their determined distances with their peculiar motion in the Galactic plane, after removing the Galactic rotation. For either models it is clear that three SFRs show a similar lag on circular rotation as the SFRs studied by Reid et al. (2009b), while two have a circular rotation close to zero lag. For NGC 281-W, the zero lag can be understood from its large distance from the Galactic disk.
Fig. 2. Phased-referenced images for two background sources belonging to maser ON 1, J2003+3034 (top) and J2009+3049 (bottom) in the third epoch. Position offset (0,0) corresponds to the position listed in Table 2. The contour levels start at a 3σ level, 6.1 and 0.6 mJy beam$^{-1}$ respectively, and increase by factors of $\sqrt{2}$.

Fig. 4. Variation of the separation between background sources J2003+3034–J2009+3049 belonging to ON 1. The solid line and the dots represent the right ascension data, while the dashed line and the filled triangles represent the declination data.

For S 255, it is not clear why its circular velocity is close to Galactic rotation (Sect. 5.6). The motion towards the Galactic Center, $U$, tends to be positive for most of the SFRs.

Recently, McMillan & Binney (2009) published a reanalysis of the maser astrometry presented by Reid et al. (2009b). They fitted the data with the revised solar peculiar motion of Binney (2009), mentioned in McMillan & Binney (2009), where the $V$ component of the solar peculiar velocity, $V_\odot$, is increased from 5.2 km s$^{-1}$ to 11 km s$^{-1}$. With a larger $V_\odot$ the lag of SFRs on Galactic rotation would decrease to ~ 11 km s$^{-1}$, which agrees with the expected velocity dispersion for young stars of ~ 10 km s$^{-1}$ (Aumer & Binney 2009). However, even when considering the revised $V_\odot$, we note that most SFRs still rotate
slower than Galactic rotation (~11 km s\(^{-1}\)) and that their velocities are not randomly dispersed.

Fig. 6. Phased-referenced images for two background sources belonging to maser L 1206, J2223+6249 (top) and J2225+6411 (bottom) in the fifth epoch. Position offset (0,0) corresponds to the position listed in Table 2. Contour levels start at a 3\(\sigma\) level, 3.5 and 0.5 mJy beam\(^{-1}\) respectively, and increase by factors of \(\sqrt{2}\).

Fig. 8. Variation of the separation between background sources J2223+6249–J2225+6411 belonging to L 1206. The solid line and the dots represent the right ascension data, while the dashed line and the filled triangles represent the declination data.

Fig. 9. Velocity integrated map and the spectrum of L.1287. Position offset (0,0) corresponds to the position listed in Table 2. Maser spots are indicated with color codes for different radial velocities. The areas of the colored circles and the colored histogram entries in the spectrum are scaled to the peak flux of that spot. Note that the black line in the spectrum is the intensity of the maser within a selected surface, which is not necessary the same as the intensity of the maser spot retrieved from a Gaussian fit. Contour levels start at 0.05 Jy beam\(^{-1}\) km s\(^{-1}\) and increase by factors of \(\sqrt{2}\).

5.2. Onsala 1

Our trigonometric parallax measurement places ON 1 on a distance from the Sun of 2.57\(^{+0.34}_{-0.27}\) kpc, somewhat closer than the (near) kinematic distance of 3.0 kpc, based on the methanol maser line at 15 km s\(^{-1}\). The latter assumes a flat rotation...
Fig. 10. Phased-referenced images for two background sources belonging to maser L 1287, J0035+6130 (top) and J0037+6236 (bottom) in epoch three. Position offset (0,0) corresponds to the position listed in Table 2. Contour levels start at a $3\sigma$ level, 3.6 and 2.4 mJy beam$^{-1}$ respectively, and increase by factors of $\sqrt{2}$.

Fig. 12. Variation of the separation between background sources J0035+6130–J0037+6236 belonging to L 1287. The solid line and the dots represent the right ascension data, while the dashed line and the filled triangles represent the declination data.

However, it is further than the commonly adopted (near) kinematic distance of 1.8 kpc based on formaldehyde at $V_{\text{LSR}} = 11.2$ km s$^{-1}$ (MacLeod et al. 1998) using the Galactic rotation curve of Wouterloot & Brand (1989), with $R_0 = 8.5$ kpc, $\Theta_0 = 220$ km s$^{-1}$. Our result resolves the near/far kinematic distance ambiguity and locates ON 1 in the Local spur consistent with the near kinematic distance.

We found different proper motions for the northern and southern maser groups (Table 5). The resultant difference in the east-west direction was small, $\mu_\alpha = 0.18 \pm 0.24$ mas yr$^{-1}$, but in the north-south direction, the southern group was moving by $\mu_\delta = -0.77 \pm 0.12$ mas yr$^{-1}$ away from the northern group. At a distance of 2.6 kpc, this corresponds to a relative speed of 9.4 km s$^{-1}$.

This large proper motion may be explained if the masers are located in the molecular gas surrounding an expanding H$\alpha$ region, as suggested by Fish & Reid (2007) and Su et al. (2009). The H$\alpha$ region is located at $(\alpha = 20^h10^m09.03, \delta = $
Fig. 14. Phased-referenced images for two background sources belonging to maser NGC 281-W, J0047+5657 (top) and J0052+5703 (bottom) in epoch five. Position offset (0,0) corresponds to the position listed in Table 2. Contour levels start at a 3σ level, 2.8 and 0.6 mJy beam$^{-1}$ respectively, and increase by factors of $\sqrt{2}$. The bottom image shows also the additional component $\sim 25$ mas westward of J0052+5703. +31°31′35.″4, J2000) and between the two maser groups. The radial velocity of the HII region, from the H76α recombination line, is $5.1 \pm 2.5$ km s$^{-1}$ (Zheng et al. 1985). The methanol masers would be in a blue-shifted component with respect to the rest frame of ON 1 at $5.1$ km s$^{-1}$, at $\sim 0$ km s$^{-1}$, northward and a red-shifted component, at $\sim 15$ km s$^{-1}$, southward of the HII region. Also, the hydroxyl masers in ON 1 provide some confirmation of this scenario (Nammahachak et al. 2006, Fish & Reid 2007). The expansion velocities of the masers, calculated by assuming a linear expansion from the center of the HII region, would be $5.8$ km s$^{-1}$ and $3.6$ km s$^{-1}$ for, respectively, the northern and southern maser groups.

5.3. L 1206

For L 1206, the trigonometric parallax distance, $0.776^{+0.11}_{-0.08}$ kpc, is smaller than the kinematic distance; 1.0 kpc (based on a hydroxyl maser, at $V_{LSR} = -8.5$ km s$^{-1}$, using the Galactic rotation curve of Wouterloot & Brand (1989) with $R_0 = 8.5$ kpc, $\Theta_0 = 220$ km s$^{-1}$, MacLeod et al. 1998), 1.23 kpc (ammonia, $V_{LSR} = -9.9$ km s$^{-1}$, using the Galactic rotation curve of Brand & Blitz (1993), with $R_0 = 8.5$ kpc, $\Theta_0 = 220$ km s$^{-1}$ Molinari et al. 1996) and 1.4 kpc (6.7 GHz methanol, flat rotation curve, $R_0 = 8.5$ kpc, $\Theta_0 = 220$ km s$^{-1}$). Our result places L 1206 in the Local spur. Cep A, a nearby (3.8 separation, P.A. 126°) star forming region, is located at a distance of $700 \pm 39$ pc, as determined by a parallax measurement of 12.2 GHz methanol masers (Moscadelli et al. 2009). The L 1206 SFR

Fig. 17. Phase-referenced image for maser S 255 at channel 4.6 km s$^{-1}$. Position offset (0,0) corresponds to the position listed in Table 2. Contour levels start at 0.23 Jy beam$^{-1}$, and increase by factors of $\sqrt{2}$.

Fig. 18. Phase-referenced images for background source J0613+1708 belonging to maser S 255, in the third epoch. Position offset (0,0) corresponds to the position listed in Table 2. The contour levels start at a 3σ level, 3.0 mJy beam$^{-1}$, and increase by factors of $\sqrt{2}$. The bottom image shows also the additional component $\sim 25$ mas westward of J0052+5703.
therefore seems to be in the same part of the Local arm as Cep A.

L 1206 is a dark cloud with an infra-red source, IRAS 22272+6358A, which coincides with the methanol maser emission and a 2.7 mm dust continuum peak (Beltrán et al. 2006).

Since neither 2 cm nor 6 cm radio emission is detected (Wilking et al. 1989; McCutcheon et al. 1991), L 1206 is thought to be in a young phase prior to the formation of an H ii region. Beltrán et al. (2006) report large CO outflows, and put the systemic and a 2.7 mm dust continuum peak (Beltrán et al. 2006).

The parallax for L 1287 sets it at a distance of 0.929±0.034 kpc from Yang et al. [1993], which placed L 1287 in the Local arm. However, the kinematic distance based on the methanol maser line would place L 1287 at 2 kpc in the Perseus arm.

Methanol masers in the dark cloud L 1287 are located at the base of the bipolar CO outflow (Yang et al. 1993) originating from the main core of the cloud. The infra-red point source (IRAS 00338+6312) in the center of the core indicates a protostellar object surrounded by cold high density NH$_3$ (1,1) gas (Estalella et al. 1993).

5.5. NGC 281-W

There are numerous distance estimates for NGC 281-W in the literature; optical photometry places the cloud between 2 and 3.68 kpc (Sharpless 1954; Cruz-González et al. 1974; Henning et al. 1994; Guetter & Turner 1977), while the kinematic distance is 3 kpc (Lee & Jung 2003), based on a $V_{LSR} = -30$ km s$^{-1}$, using the Galactic rotation curve of Clemens (1985) and $R_0 = 8.5$ kpc, $\theta_0 = 220$ km s$^{-1}$. The radial velocity of the methanol maser line at $-29$ km s$^{-1}$, assuming a flat Galactic rotation curve, and $R_0 = 8.5$ kpc, $\theta_0 = 220$ km s$^{-1}$, would place NGC 281-W at 2.5 kpc.

Recently, Sato et al. (2008) reported a water maser parallax value of 0.355 ± 0.030 mas, measured with the VERA interferometer, corresponding to a distance of 2.82 ± 0.24 kpc. We measured a parallax of 0.421 ± 0.022 mas, arriving at a distance of 2.38$^{+0.13}_{-0.12}$ kpc. These distance measurements agree with Sato et al. (2008) within 2$\sigma$ of the joint uncertainty. Also, our proper motion agrees within 1$\sigma$ with the proper motion by Sato et al. (2008), and our radial velocity ($V_{LSR} \sim -29$ km s$^{-1}$) is close to that of the water masers ($V_{LSR} \sim -31$ km s$^{-1}$). It is unlikely that the water and methanol masers are originated at such different depths into the SFR that it would affect the distance. For example, in the SFR W3(OH) the water and methanol maser parallaxes are in good agreement (Xu et al. 2006; Hachisuka et al. 2006).

NGC 281-W is a SFR that lies near the edge of a super bubble in the Perseus arm (Megerle et al. 2003; Sato et al. 2007). Our slightly smaller distance would change the position of NGC 281-W on this super bubble. Another SFR located on this bubble has a parallax determined distance: IRAS 00420+5530 at 2.17 ± 0.05 kpc (Moellenbrock et al. 2009). This allows to compare the distances and peculiar motions of both SFRs, and...
check if the expanding super bubble, with the expansion center at 2.5 kpc adopted after Sato et al. (2008), can be responsible for the peculiar motions of both of them. We calculated the peculiar motions of IRAS 00420+5530 after Moellenbrock et al. (2009) and NGC 281-W after Sato et al. (2008) and our results. Fig. 21 bottom panel, shows the peculiar motions in a face-on view of the Galactic plane. From the figure can be seen that the $U$ component of NGC 281-W, representing the vector toward the Galactic center, is positive and non-zero in both our study and the work of Sato et al. (2008). This means that NGC 281-W is moving toward the Galactic center, which is the expected movement for the near side of an expanding bubble. In the top panel of Fig. 21 the peculiar motions of both SFRs are shown projected on a cross-section of the Galactic plane along the longitude of NGC 281-W, $l = 123^\circ$. The average distance to NGC 281-W based on Sato et al. (2008) and this work, 2.58±0.23 kpc, is marked in this panel as a cross. If we prolong the peculiar motion backwards of both IRAS 00420+5530 and the average result for NGC 281-W, they intersect at a distance from the Sun of 2.65 kpc and at a Galactic latitude of $-0.1^\circ$. This suggests that if indeed both SFRs are expanding linearly from one expansion center, that the expansion center is offset from the previously assumed position. However, this offset center of expansion is difficult to understand when looking to the peculiar motion of both SFRs in the Galactic plane (Fig. 21 lower panel), since the motion does not seem to originate from one mutual expansion center. Hence, it is difficult to pinpoint the characteristics of the super bubble, with peculiar motions and distances of only two SFRs known.

6. Summary

We have measured parallaxes of 6.7 GHz methanol masers using the EVN towards five SFRs, achieving accuracies approaching 20 μas. The primary results are summarized as follows:

1. We report trigonometric parallaxes for five star forming regions, the distances to these sources are 2.57±0.37 kpc for ON 1, 0.776±0.104 kpc for L 1206, 0.929±0.034 kpc for L 1287, 2.38±0.13 kpc for NGC 281-W and 1.59±0.06 kpc for S 255.

2. Galactic star forming regions lag circular rotation on average by $-17$ km s$^{-1}$, a value comparable to that found recently by similar studies (Reid et al. 2009).

3. Individual 6.7 GHz methanol maser spots are stable over a period of > 2 years for most of the maser spots. The relative motions between the maser spots are small ($\sim 0.5–1$ km s$^{-1}$) and exhibit a rectilinear motion.

4. Measurements at 6.7 GHz are less disturbed from the troposphere, as expected. However, the ionospheric delay cannot be ignored and is likely not completely removed by using the JPL GPS–IONEX maps to calculate propagation delays. Continuum measurements of background sources show that most of them have significant structure evident from the large (up to 1–2 mas, or 2–10 km s$^{-1}$) apparent movements between pairs of background sources. This additional uncertainty from the apparent movement does not prohibit a determination of the peculiar motion, however, it increases the error bars of the peculiar motion by the order of the uncertainty, depending on the distance and longitude of the source. The parallax uncertainty can be affected as well, if the cause of the apparent movement between the background sources are internal structure changes, which do not need to be perfectly linear. Two background sources should be a minimum for astrometric measurements at 6.7 GHz and three or more background sources are recommended.

Acknowledgements. The European VLBI Network is a joint facility of European, Chinese, South African and other radio astronomy institutes funded...
The website contains several references to scientific papers and authors, but the content is not clearly visible due to the image quality. The text appears to be a mix of scientific publications and possibly conference proceedings, with authors such as Reid, Menten, Brunthaler, and others. The references cover a range of years from 2002 to 2009, indicating that the content is likely from recent research in astronomy.

For a more detailed analysis, the text would need to be readable, which is not the case here. If you have a clearer version of the text, I would be able to provide a more comprehensive analysis.
Table 5. Detailed results of parallax and proper motions measurements

| Background source | $V_{\text{LSR}}$ (km s$^{-1}$) | Parallax (mas) | $\mu_\alpha$ (mas yr$^{-1}$) | $\mu_\delta$ (mas yr$^{-1}$) |
|-------------------|-------------------------------|----------------|-----------------------------|-----------------------------|
| **ON 1**          |                               |                |                             |                             |
| **Northern group**|                               |                |                             |                             |
| J2003+3034        | -0.4                          | 0.521 ± 0.055  | -3.33 ± 0.18                | -5.26 ± 0.41                |
|                   | 0.0                           | 0.401 ± 0.052  | -3.72 ± 0.17                | -5.26 ± 0.21                |
|                   | 0.4                           | 0.361 ± 0.052  | -3.82 ± 0.17                | -5.28 ± 0.19                |
|                   | 0.7                           | 0.526 ± 0.057  | -3.36 ± 0.19                | -5.19 ± 0.34                |
| Combined fit      |                               | 0.299 ± 0.112$^1$ |                             |                             |
| Averaging data    |                               | 0.391 ± 0.061  |                             |                             |
| J2009+3049        | -0.4                          | 0.380 ± 0.121  | -2.87 ± 0.39                | -4.88 ± 0.31                |
|                   | 0.0                           | 0.459 ± 0.140  | -2.68 ± 0.46                | -4.80 ± 0.22                |
|                   | 0.4                           | 0.483 ± 0.142  | -2.59 ± 0.48                | -4.79 ± 0.20                |
|                   | 0.7                           | 0.415 ± 0.112  | -2.81 ± 0.37                | -4.78 ± 0.23                |
| Combined fit      |                               | 0.482 ± 0.137$^1$ |                             |                             |
| Averaging data    |                               | 0.368 ± 0.070  |                             |                             |
| **Southern group**|                               |                |                             |                             |
| J2003+3034        | 14.4                          | -3.88 ± 0.12   | -5.92 ± 0.14                |                             |
|                   | 14.8                          | -3.75 ± 0.09   | -6.03 ± 0.11                |                             |
|                   | 15.1                          | -3.67 ± 0.21   | -6.09 ± 0.19                |                             |
| J2009+3049        | 14.4                          | -3.00 ± 0.15   | -5.49 ± 0.20                |                             |
|                   | 14.8                          | -2.88 ± 0.21   | -5.59 ± 0.26                |                             |
|                   | 15.1                          | -2.77 ± 0.29   | -5.65 ± 0.35                |                             |
| Both QSOs         |                               | 0.390 ± 0.116$^1$ |                             |                             |
| Averaging data    |                               | 0.389 ± 0.045  |                             |                             |

$<\mu>_{\text{north}}$ (σ)$_{pm}$

$<\mu>_{\text{south}}$ (σ)$_{pm}$

1. This is the error of the combined fit multiplied by $\sqrt{N}$, where $N$ is the number of maser spots.
2. We calculated an unweighted arithmetic mean of the individual proper motion results from all maser spots and background sources. The error bar on the mean is the standard error of the mean to which was added, in quadrature, the apparent movement between the two background sources of the respective coordinate. The uncertainty in the proper motion, which is introduced by the background source, is hereby taken into account. In parenthesis is given the standard deviation of the mean.
Table 5. continued.

| Background source | $V_{LSR}$ (km s$^{-1}$) | Parallax (mas) | $\mu_\alpha$ (mas yr$^{-1}$) | $\mu_\delta$ (mas yr$^{-1}$) |
|-------------------|-------------------------|----------------|-----------------------------|-------------------------------|
| Combined fit      | $0.984 \pm 0.086^1$     | $-0.14 \pm 0.10$ | $-1.71 \pm 0.40$            |
| Averaging data    | $1.016 \pm 0.052$       |                |                             |
| J0037+6236        | $-27.0$                 | $1.306 \pm 0.071$ | $-0.88 \pm 0.04$            | $-1.72 \pm 0.27$             |
|                   | $-23.9$                 | $1.244 \pm 0.044$ | $-0.81 \pm 0.04$            | $-1.77 \pm 0.25$             |
|                   | $-23.5$                 | $1.225 \pm 0.040$ | $-0.88 \pm 0.14$            | $-1.83 \pm 0.12$             |
|                   | $-23.2$                 | $1.011 \pm 0.101$ | $-1.17 \pm 0.20$            | $-2.20 \pm 0.14$             |
|                   | $-22.5$                 | $1.330 \pm 0.043$ | $-1.01 \pm 0.04$            | $-2.64 \pm 0.37$             |
| Combined fit      | $1.192 \pm 0.107^1$     |                |                             |
| Averaging data    | $1.150 \pm 0.052$       |                |                             |
| Both QSOs         | Combined fit            | $1.079 \pm 0.069^1$ | $-0.86 \pm 0.11(0.33)^2$   | $-2.29 \pm 0.56(0.46)^2$ |
| Averaging data    | $1.077 \pm 0.039$       |                |                             |
| $< \mu > (\sigma)_\mu$ |                  | $-2.64 \pm 0.15$ | $-1.75 \pm 0.10$            |
| NGC 281-W         |                         |                |                             |
| J0047+5657        | $-30.2$                 | $0.416 \pm 0.054$ | $-2.64 \pm 0.15$            | $-1.75 \pm 0.10$             |
|                   | $-29.9$                 | $0.380 \pm 0.045$ | $-2.54 \pm 0.12$            | $-1.72 \pm 0.08$             |
|                   | $-29.5$                 | $0.401 \pm 0.041$ | $-2.58 \pm 0.19$            | $-1.72 \pm 0.07$             |
|                   | $-29.2$                 | $0.404 \pm 0.047$ | $-2.69 \pm 0.22$            | $-1.75 \pm 0.08$             |
|                   | $-28.8$                 | $0.497 \pm 0.026$ | $-2.55 \pm 0.04$            | $-1.69 \pm 0.13$             |
|                   | $-28.1$                 | $0.529 \pm 0.028$ | $-2.68 \pm 0.04$            | $-1.52 \pm 0.14$             |
| Combined fit      | $0.400 \pm 0.070^1$     |                |                             |
| Averaging data    | $0.398 \pm 0.042$       |                |                             |
| J0052+5703        | $-30.2$                 | $0.459 \pm 0.049$ | $-2.79 \pm 0.06$            | $-1.87 \pm 0.27$             |
|                   | $-29.9$                 | $0.420 \pm 0.048$ | $-2.68 \pm 0.05$            | $-1.84 \pm 0.25$             |
|                   | $-29.5$                 | $0.388 \pm 0.048$ | $-2.75 \pm 0.05$            | $-1.84 \pm 0.22$             |
|                   | $-29.2$                 | $0.408 \pm 0.054^1$ | $-2.87 \pm 0.07$            | $-1.87 \pm 0.20$             |
|                   | $-28.8$                 | $0.602 \pm 0.056$ | $-2.65 \pm 0.07$            | $-1.82 \pm 0.26$             |
|                   | $-28.1$                 | $0.648 \pm 0.055$ | $-2.76 \pm 0.07$            | $-1.65 \pm 0.34$             |
| Combined fit      | $0.399 \pm 0.054^1$     |                |                             |
| Averaging data    | $0.425 \pm 0.024$       |                |                             |
| Both QSOs         | Combined fit            | $0.412 \pm 0.045^1$ | $-2.69 \pm 0.16(0.10)^2$   | $-1.77 \pm 0.11(0.10)^2$ |
| Averaging data    | $0.421 \pm 0.022$       |                |                             |
| $< \mu > (\sigma)_\mu$ |                  | $-2.64 \pm 0.15$ | $-1.75 \pm 0.10$            |
| S 255             |                         |                |                             |
| J0613+1708        | $4.6$                   | $0.628 \pm 0.027$ | $-0.14 \pm 0.05$            | $-0.84 \pm 1.67$             |
| $\mu_\mu$        |                         | $-0.14 \pm 0.54^2$ | $-0.84 \pm 1.76^2$          |