The brightest Ly α emitter: Pop III or black hole?

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ABSTRACT

CR7 is the brightest $z = 6.6$ Ly α emitter (LAE) known to date, and spectroscopic follow-up by Sobral et al. suggests that CR7 might host Population (Pop) III stars. We examine this interpretation using cosmological hydrodynamical simulations. Several simulated galaxies show the same ‘Pop III wave’ pattern observed in CR7. However, to reproduce the extreme CR7 Ly α/He ii line luminosities ($L_{\alpha}/L_{\text{He} \text{II}}$) a top-heavy initial mass function and a massive ($\gtrsim 10^7 M_\odot$) Pop III burst with age $\lesssim 2$ Myr are required. Assuming that the observed properties of Ly α and He ii emission are typical for Pop III, we predict that in the COSMOS/UDS/SA22 fields, 14 out of the 30 LAEs at $z = 6.6$ with $L_\alpha > 10^{43.3} \text{ erg s}^{-1}$ should also host Pop III stars producing an observable $L_{\text{He} \text{II}} \gtrsim 10^{42.7} \text{ erg s}^{-1}$. As an alternate explanation, we explore the possibility that CR7 is instead powered by accretion on to a direct collapse black hole. Our model predicts $L_\alpha$, $L_{\text{He} \text{II}}$, and X-ray luminosities that are in agreement with the observations. In any case, the observed properties of CR7 indicate that this galaxy is most likely powered by sources formed from pristine gas. We propose that further X-ray observations can distinguish between the two above scenarios.

Key words: black hole physics – stars: Population III – galaxies: high-redshift.

1 INTRODUCTION

The end of the dark ages is marked by the appearance of the first stars. Such – Population (Pop) III – stars had to form out of a pristine composition (H+He) gas with virtually no heavy elements. Lacking these cooling agents, the collapse had to rely on the inefficient radiative losses provided by H$_2$ molecules. Mini-haloes, i.e. non-linear dark matter structures with mass $M_\text{h} \sim 10^{6-7} M_\odot$ collapsing at high redshift ($z \sim 30$), are now thought to be the preferred sites of first star formation episodes (Yoshida et al. 2006; Salvadori & Ferrara 2009; Turk, Abel & O’Shea 2009; Greif et al. 2012; Visbal, Haiman & Bryan 2015). Although the initial mass function (IMF) of Pop III stars is largely uncertain, physical arguments suggest that they could have been more massive than present-day (Pop II) stars. Furthermore, the metals produced by Pop III stars polluted the surrounding gas (Bromm, Coppi & Larson 2002; Wise et al. 2012; Xu, Wise & Norman 2013), inducing a transition to the Pop II star formation mode (‘chemical feedback’; Schneider et al. 2002, 2006). Metal enrichment is far from being homogeneous, and pockets of pristine gas sustaining Pop III star formation can in principle persist down to $z \sim 3-4$ (Tornatore, Ferrara & Schneider 2007; Trenti, Stiavelli & Michael Shull 2009; Maio et al. 2010; Pallottini et al. 2014a; Salvadori et al. 2014; Ma et al. 2015), yielding Pop III star formation rate (SFR) densities of $\sim 10^{-2} M_\odot \text{ yr}^{-1} \text{ Mpc}^{-3}$, i.e. $\lesssim 1$ per cent of the Pop II SFR density at those redshifts.

The search effort for Pop III stars at moderate and high redshifts has become increasingly intense in the last few years (e.g. Kashikawa et al. 2012; Heap, Bouret & Hubeny 2015). Observationally, a galaxy hosting a recent ($t_{\text{popIII}} \lesssim 2$ Myr) Pop III star formation episode should show strong Lyα and He ii lines and no metal lines (e.g. Schaerer 2002; Raiter, Schaerer & Fosbury 2010; Kehrig et al. 2015). Until now, no indisputable evidence for Pop III stars in distant galaxies has been obtained, and observations have only yielded upper bounds on Pop III SFR (e.g. Cai et al. 2011; Cassata et al. 2013; Zabl et al. 2015). This situation might dramatically change...
following the recent observations of CR7 by Sobral et al. (2015, S15 hereafter).

CR7 is the brightest Ly α emitter (LAE) at $z > 6$, and it is found in the COSMOS field (Matthee et al. 2015). Spectroscopic follow-up by S15 suggests that CR7 might host a Pop III-like stellar population. This is based on the astonishingly bright Ly α and He ii lines ($L_{\text{Ly} \alpha} \simeq 10^{41.93} \text{erg s}^{-1}$, $L_{\text{He} \text{II}} \simeq 10^{41.20} \text{erg s}^{-1}$) and no detection of metal lines. S15 show that CR7 can be described by a composite of a Pop III-like and a more normal stellar population, which would have to be physically separated, and that would be consistent with e.g. Tornatore et al. (2007). Hubble Space Telescope imaging shows that CR7 is indeed composed of different components: three separate sub-systems (A, B, C) with projected separations of $\lesssim 5$ kpc. $F110W(I\gamma)$ and $F160W(H\beta)$-band photometry indicates that clump A might be composed of young (blue) stars, while the stellar populations of B+C are old and relatively red. The observed Ly α and He ii lines are narrow (FWHM $\lesssim 200$ km s$^{-1}$ and FWHM $\lesssim 130$ km s$^{-1}$, respectively), disfavouring the presence of an active galactic nucleus or Wolf–Rayet stars, which are expected to produce much broader (FWHM $\gtrsim 10^5$ km s$^{-1}$) lines (e.g. De Breuck et al. 2000; Brinchmann, Pettini & Charlot 2008; Erb et al. 2010). S15 concluded that CR7 likely contains a composite stellar population, with clump A being powered by a recent Pop III-like burst ($t_1 \lesssim 2$ Myr), and clumps B+C containing an old ($t_1 \sim 350$ Myr) burst of Pop II stars with $M_* \simeq 10^{10}$ M$\odot$, largely dominating the stellar mass of the entire system.

Based on cosmological simulations that follows the simultaneous evolution of Pop II and Pop III stars (Pallottini et al. 2014a, P14 hereafter), we examine the interpretation of CR7 as a Pop III host system and explore its implications. We also propose an alternate explanation, briefly discussed in S15, where CR7 is powered by accretion on to a direct collapse black hole (DCBH) and suggest further tests.

2 SIMULATION OVERVIEW

We use the $\Lambda$ cold dark matter ($\Lambda$CDM) cosmological$^1$ hydrodynamical simulations presented in P14 (see that paper for a comprehensive description), obtained with a customized version of the adaptive mesh refinement code RAMSES (Teyssier 2002) to evolve a $(10 h^{-1}$ Mpc)$^3$ volume from $z = 99$ to 4, with a dark matter mass resolution of $\simeq 5 \times 10^5 M_{\odot}$, and an adaptive baryon spatial resolution ranging from $\simeq 20$ to $\simeq 1 h^{-1}$ kpc. Star formation is included via sub-grid prescriptions based on a local density threshold. If the star-forming cell gas has metallicity below (above) the critical metallicity, $Z_{\text{crit}} = 10^{-3} Z_{\odot}$, we label the newly formed stars as Pop III (Pop II). Supernova feedback accounts for metal-dependent stellar yields and return fractions appropriate for the relevant stellar population.$^2$ The simulated galaxy sample reproduces the observed cosmic SFR (Bouwens et al. 2012; Zheng et al. 2012) and stellar mass density (González et al. 2011) evolution in the redshift range 4 $\lesssim z \lesssim 10$, and as shown in Pallottini et al. (2015) – P14 reproduce the observed luminosity function at $z = 6$. Additionally, the derived Pop III cosmic SFR density is consistent with current observational upper limits (e.g. Nagao et al. 2008; Cai et al. 2011; Cassata et al. 2013). To allow a direct comparison with CR7, we will concentrate on the analysis of the $z \simeq 6$ simulation output.

2.1 Pop III-hosting galaxies

As noted in S15, the interpretation of CR7 fits in the ‘Pop III wave’ scenario suggested by Tornatore et al. (2007). As an example of Pop III wave in action, we show the case of ‘MB45’, a simulated P14 galaxy with total stellar mass $M_* = 10^{10.9}$ M$\odot$. In Fig. 1, we plot the metallicity ($Z$) and overdensity ($\Delta$) map around MB45. The star formation history in MB45 starts with a Pop III event. These stars explode as supernovae enriching with metals the central regions of MB45. As a result, star formation there continues in the Pop II mode, while in the less-dense external regions, not yet reached by the metal-bearing shocks, Pop III stars can still form. The process repeats until the unpolluted regions have densities exceedingly low to sustain star formation.

The total (i.e. old+young stars) Pop III mass in MB45 is $M_1 \simeq 10^{8.6}$ M$\odot$; about 20 per cent of this stellar mass formed in a recent burst ($age t_1 \lesssim 2$ Myr). The total stellar mass ($M_* \simeq 10^8$ M$\odot$) of MB45 is dominated by Pop II stars produced at a rate $\text{SFR}_2 \simeq 0.5 M_{\odot}$ yr$^{-1}$. Thus, while MB45 formation activity proceeds in the Pop-III-wave mode and resembles that of CR7, the physical properties of MB45 and CR7 are different, because of the two orders of magnitude difference in total stellar mass. A direct comparison between the Pop III–Pop II separation in CR7 (projected $\lesssim 5$ kpc) and MB45 (10 kpc, projected $\lesssim 5$ kpc), although fairly consistent, might not be very meaningful due to the different mass of the two systems. This is because the separation depends on the mass-dependent metallicity profile in galaxy groups (see fig. 1 in Pallottini, Galleroni & Ferrara 2014b). However, we note that our current simulated volume is simply too small to be able to directly recover sources such as CR7, with volume densities of $\sim 10^{-6}$ Mpc$^{-3}$.

$^1$ We assume a $\Lambda$CDM cosmology with vacuum, total matter, and baryonic densities in units of the critical density $\Omega_{\Lambda} = 0.727$, $\Omega_{\text{dm}} = 0.228$, $\Omega_b = 0.045$, Hubble constant $H_0 = 100 h \text{ km s}^{-1} \text{ Mpc}^{-1}$ with $h = 0.704$, spectral index $n = 0.967$, and $\sigma_8 = 0.811$ ( Larson et al. 2011).

$^2$ While in P14 we explore different types of IMF for Pop III, here we show results assuming a Pop II-like Salpeter IMF. It is to note that our simulations suggest that the Pop III SFR seems almost independent from the IMF (see in particular fig. 14 in P14).
no He II emission. For a composite galaxy, the Ly α emission is the sum of the contributions from Pop II and Pop III components:

\[ L_\alpha^{\text{comp}} = L_\alpha^{\text{pure}} + A_\alpha l_\alpha^3 M_3, \]  

(1b)

where \( l_\alpha \) is the Pop III Ly α line luminosity per unit stellar mass.\(^4\) Analogously, the He II emission is given by

\[ L_{\text{HeII}} = l_{\text{HeII}}^3 M_3, \]  

(1c)

where \( l_{\text{HeII}} \) is the Pop III He II luminosity per unit stellar mass. Both \( l_\alpha \) and \( l_{\text{HeII}} \) depend on the IMF and burst age, \( t_b \). We adopt the Pop III models by Schauer (2002) and Raitter et al. (2010), and we use a Salpeter IMF (power-law slope \( \alpha = -2.35 \)), with variable lower \( (m_{\text{low}}) \) and upper \( (m_{\text{up}}) \) limits. As long as \( m_{\text{up}} \gtrsim 10^5 M_\odot \), the results are very weakly dependent on the upper limit, which we therefore fix to \( m_{\text{up}} = 10^6 M_\odot \), leaving \( m_{\text{low}} \) as the only free parameter.

As noted by S15, to reproduce CR7 \( L_\alpha \) and \( L_{\text{HeII}} \) with Pop III stars, a mass of \( M_3 \simeq 10^{10} M_\odot \) newly-born \((t_b \lesssim 2-5 \text{ Myr})\) stars is required, depending on the IMF. Such a large amount of young Pop III stars is contained in none of the P14 galaxies and it is not predicted by the adopted analytical extrapolation (see Fig. 2). Thus, none of the simulated composite galaxies would reproduce CR7 line emission. However, it is possible that CR7 might have experienced a more vigorous Pop III star formation burst as a result of a very rare event – e.g. a recent major merger – not frequent enough to be captured in our limited box volume. As an estimate, we adopt the value of \( M_3 \) resulting from the sum of all (old+young) Pop III stars formed in our galaxies.

Under this hypothesis, we can fix \( m_{\text{low}} \), by using the zero-age main-sequence (ZAMS) tracks \((t_b = 0)\). By SED fitting, S15 shows that CR7 Pop II stellar mass (completely contained in clumps B+C) is likely \( M_3 \simeq 5 \times 10^{10} M_\odot \). From the lower panel of Fig. 2, we find that this mass corresponds to a Pop III mass of \( M_3 \sim 10^{7.5} M_\odot \). From Pop III SED fits of region A S15 estimate \( M_3 \sim 10^{7} M_\odot \). As these stars must be located in CR7 clump A, whose He II luminosity is \( L_{\text{HeII}} = 10^{43.3} \text{ erg s}^{-1} \), equation (1c) requires that \( l_{\text{HeII}} = 10^{54.5} \text{ erg s}^{-1} M_\odot^{-1} \). In turn this entails a top-heavy IMF with \( m_{\text{low}} = 6.7 M_\odot \).

Having fixed the IMF, we can readily derive the predicted Pop III contribution to the Ly α emission; this turns out to be \( l_\alpha^3 = 10^{66.7} \text{ erg s}^{-1} M_\odot^{-3} \). CR7 has an observed \( L_\alpha = 10^{43.9} \text{ erg s}^{-1} \), with no contribution from clumps B+C (SFR2 = 0). This comparison allows us to determine, using equation (1b), the Ly α line attenuation factor, \( A_\alpha = 10^{-0.57} \).

Roughly 66 per cent of the line luminosity is therefore damped, a figure consistent with other derivations \((\text{e.g. Dayal et al. 2008})\), and with the analysis of S15. The above procedure provides a basis to model Ly α and He II emission for both pure and composite galaxies, assuming that the properties of the Pop III component are similar to those derived from CR7.

### 3 Predictions for bright LAEs

Starting from the assumption that CR7 is a ‘typical’ composite galaxy, and using \( m_{\text{low}} = 6.7 M_\odot \) and \( A_\alpha = 10^{-0.57} \), we can now predict how many LAEs among those observed by Matthee et al. (2015) are composite galaxies, i.e. contain Pop III stars. The number of LAEs in the COSMOS/UDS/SA22 fields with luminosity \( L_\alpha \) is \( N_\alpha = \Phi_\alpha(L_\alpha) V_{\text{obs}} \), where \( \Phi_\alpha \) is the observed Ly α luminosity.

\(^3\) We refer to appendix A of P14 for possible resolution effects.

\(^4\) We are implicitly assuming that Pop III stars form in a burst, an assumption justified by the analysis presented in Section 2.1.
function and $V_{\text{obs}} = 4.26 \times 10^9 \text{Mpc}^3$ is the observed volume. Among these, a fraction

$$f_{\alpha}^{\text{comp}} = N_{\text{comp}} / (N_{\text{comp}} + N_{\text{pure}})$$  \hspace{1cm} (2a)$$

contains Pop III stars, where $N_{\text{comp}}$ ($N_{\text{pure}}$) is the number of composite (pure) galaxies in $V_{\text{obs}}$ at a given $L_{\alpha}$. Equations (1a) and (1b), show that a given $L_{\alpha}$ can be produced by a composite galaxy with a lower $M_{\star}$ with respect to a pure (Pop II) galaxy. For instance, $L_{\alpha} \approx 10^{43.5}$ erg s$^{-1}$ requires $M_{\star} \approx 10^{10.5} M_{\odot}$ for a pure galaxy, but only $M_{\star} \lesssim 10^{10} M_{\odot}$ for a composite one. Such large objects are very rare at the redshift of CR7 ($z = 6.6$) and in the observed volume.\footnote{As a reference, an $M_{\star} \sim 10^{10.5} M_{\odot}$ is hosted in a dark matter halo of mass $M_{h} \sim 10^{13.5} M_{\odot}$, whose abundance is $n_{h} \sim 10^{-8} \text{Mpc}^{-3}$ at $z \approx 6$ (e.g. Sheth & Tormen 1999).} Therefore, it is important to account for the statistical (Poisson) fluctuations of the galaxy number counts as shown in the upper panel of Fig. 2 as the 'young' Pop III curve,\footnote{As noted in Section 2.1, all galaxies with $M_{\star} > 10^7$ have old Pop III stars, thus considering the 'old-young' track for $f_{\text{comp}}$ would yield an unrealistically high composite number.} while $n_{\text{pure}}$ accounts for the remaining galaxies. The effect of the finite volume effects on $f_{\alpha}$\footnote{The predicted number would become $7 \pm 7$, by assuming sSFR from Stark et al. (2013).} can be appreciated from the lower panel of Fig. 3. Assuming a higher sSFR = 5 Gyr$^{-1}$ (e.g. Stark et al. 2013) yields the modifications shown by the dash–dotted line.

In the upper panel of Fig. 3, we plot the LAE number ($N_{\alpha}$) as a function of $L_{\alpha}$, Matthee et al. (2015) observations (green pentagons) are shown along with our predictions for the composite LAE and expected $L_{\text{HeII}}$ emission. CR7 is the most luminous LAE observed, and it is in the brightest luminosity bin ($L_{\alpha} = 10^{44.1 \pm 0.1}$ erg s$^{-1}$); by assumption CR7 is a composite galaxy. If so, we then predict that out of the 46 (30) LAEs with $L_{\alpha} = 10^{43.2 \pm 0.1}$ erg s$^{-1}$ ($> 10^{43.5}$ erg s$^{-1}$, cumulative), 13 (14) must also be composite galaxies,\footnote{As recomputed in S15 using the $Y$ band to estimate the continuum, in order to match the calculation for CR7.} with observable $L_{\text{HeII}} \approx 10^{42.5}$ erg s$^{-1}$ ($> 10^{42.7}$ erg s$^{-1}$). Follow-up spectroscopy of those luminous LAEs at $z = 6.6$ will allow us to test this prediction. We recall that this test assumes that all Pop III give rise to the same Ly$\alpha$ and He ii emission as inferred from CR7, that requires that all the Pop III stellar mass was formed in a single burst with age $\lesssim 2$ Myr.

Particularly in the regime where $f_{\text{comp}} < 1$ (see the lower panel in Fig. 3), a sample of LAEs is needed to test our model predictions. For example for 'Himiko', the second most luminous\footnote{The LW is estimated by accounting for the stellar properties of clump B+C (in particular see fig. 8 in S15), and by assuming a 5 kpc distance between B+C and A.} confirmed LAE at $z = 6.6$ with $L_{\alpha} \approx 10^{44.3}$ erg s$^{-1}$ (Ouchi et al. 2009), for which recent VLT/X-Shooter observations have provided a 3$\sigma$ limit of $L_{\text{HeII}} \lesssim 10^{42.2}$ erg s$^{-1}$ (Zabl et al. 2015), our model predicts $L_{\text{HeII}} \lesssim 10^{42.7}$ erg s$^{-1}$, i.e. a four times higher He ii luminosity. However, this is predicted only for $f_{\text{comp}} \approx 20–30$ per cent of galaxies at this $L_{\alpha}$.

4 ALTERNATIVE INTERPRETATION

Given the extreme conditions required to explain the observed properties of CR7 in terms of Pop III stars and a set of assumptions, it is worth exploring alternative interpretations. The most appealing one involves DCBHs, which is briefly discussed in S15. High-z pristine, atomic haloes ($M_{h} \sim 10^{8} M_{\odot}$) primarily cool via Ly$\alpha$ line emission. In the presence of an intense Lyman–Werner (LW, $E_{\nu} = 112–13.6$ eV) irradiation, $H_{2}$ molecule photodissociation enforces an isothermal collapse (Shang, Bryan & Haiman 2010; Agarwal et al. 2013; Latif et al. 2013; Yue et al. 2014), finally leading to the formation of a DCBH of initial mass $M_{\star} \approx 10^{5.3–5.5} M_{\odot}$ (Begelman, Volonteri & Rees 2006; Volonteri, Lodato & Natarajan 2008; Ferrara et al. 2014), eventually growing up to $10^{6–7} M_{\odot}$ by accretion of the halo leftover gas.

In CR7, clump A appears to be pristine, and it is irradiated by an LW flux from B+C\footnote{The LW is estimated by accounting for the stellar properties of clump B+C (in particular see fig. 8 in S15), and by assuming a 5 kpc distance between B+C and A.} of $\sim 5 \times 10^{-18}$ erg s$^{-1}$ cm$^{-2}$ Hz$^{-1}$ sr$^{-1}$, well in excess of the required threshold for DCBH formation (Shang et al. 2010; Latif et al. 2013; Regan, Johansson & Wise 2014; Sugimura, Omukai & Inoue 2014). Thus, CR7 might be a perfect host for a DCBH.

We investigate the time-evolving spectrum of an accreting DCBH of initial mass $M_{\star} = 10^{4.5} M_{\odot}$ by coupling a 1D radiation-hydrodynamic code (Pacucci & Ferrara 2015) to the spectral synthesis code CLOUDY (Ferland et al. 2013), as detailed in Pacucci et al. (2015). The DCBH intrinsic spectrum is taken from Yue et al. (2013). The DCBH is at the centre of a halo of total gas mass $M_{h} \sim 10^{8} M_{\odot}$, distributed with a core plus an $r^{-2}$ density profile spanning up to 10 pc. The accretion is followed until complete depletion of the halo gas, i.e. for $\lesssim 120$ Myr. During this period, the total absorbing column density of the gas varies from an initial value of $\lesssim 3.5 \times 10^{22}$ cm$^{-2}$ to a final value $< 10^{21}$ cm$^{-2}$, i.e. from mildly Compton-thick to strongly Compton-thin. Note that while Ly $\alpha$ attenuation by the interstellar medium is included, we do not account for the likely sub-dominant IGM analogous effect.

Fig. 4 shows the time evolution of the Ly $\alpha$, He ii, and X-ray (0.5–2 keV) luminosities. Both Ly $\alpha$ and He ii are consistent with the
observed CR7 values during an evolutionary phase lasting ~17 Myr (14 per cent of the system lifetime), longer than the shorter period (t, ≲ 2 Myr) of our assumption for a massive Pop III burst.

The equivalent width of the He II line in the CR7 compatibility region ranges from 75 to 85 Å. The column density during the CR7-compatible period is ≳ 10^25 cm^-2, i.e. mildly Compton-thick. The associated X-ray luminosity is ≳ 10^{43} erg s^{-1}, fully consistent with the current upper limit for CR7 (≲ 10^{44} erg s^{-1}; Elvis et al. 2009). Deeper X-ray observations of CR7 might then confirm the presence of the DCBH. However, this limit is already obtained with 180 ks of integration time on Chandra, meaning that a stringent test might only be possible with the next generation of X-ray telescopes.

5 CONCLUSIONS

CR7 is the brightest z = 6.6 LAE in the COSMOS field (Matthee et al. 2015). Spectroscopic follow-up (Sobral et al. 2015) suggests that CR7 might host Pop III stars, along with Pop II and thus be explained by a ‘Pop III wave’ scenario. We have further investigated such interpretation using cosmological simulations following the formation of Pop II and Pop III stars in early galaxies.

We find simulated galaxies (like MB45 in Fig. 1) hosting both Pop III and Pop II stars at z = 6.0. Such ‘composite’ galaxies have morphologies similar to that of CR7 and consistent with the ‘Pop III wave scenario’. However, to reproduce the extreme CR7 Ly α/He II 1640 line luminosities, a top-heavy IMF combined with a massive (M_\text{HI} ≳ 10^5 M_\odot) Pop III burst of young stars (t, ≲ 2−5 Myr) is required. Our simulations do not predict such large burst, i.e. M_\text{HI} ≲ 10^6 M_\odot, but our volume is also smaller than that used to discover CR7. None of the less, assuming that CR7 is typical of all metal-free components in our simulations, we predict that in the combined COSMOS, UDS and SA22 fields, out of the 30 LAEs with L_\text{Ly} > 10^{43.3} erg s^{-1}, 14 should also host Pop III stars producing an observable L_{\text{He II}} ≳ 10^{42.7} erg s^{-1}.

Given the extreme requirements set by the Pop III interpretation, we explored the possibility that CR7 is instead powered by accretion on to a DCBH of initial mass 10^5 M_\odot. The predicted L_{\text{He II}} match CR7 observations during a time interval of ~17 Myr (~14 per cent of the system lifetime). The predicted CR7 luminosity at 0.5–2 keV, ≲ 10^{43} erg s^{-1}, is significantly below the current upper limit, i.e. ≲ 10^{44} erg s^{-1}.

We conclude that the DCBH interpretation of CR7 is very appealing, and competitive with the explanation involving a massive Pop III burst. For both explanations, the dominant ionizing source of this galaxy should have formed from pristine gas. Deep X-ray observations and other follow-up observations should allow us to shed more light on this very peculiar source.

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