1. INTRODUCTION

Unlike the optical, X-ray, and γ-ray sky, which have been systematically studied in the time domain in the search for supernovae (SNe) and gamma-ray bursts (GRBs), the wide-field ultraviolet (UV) time domain is a relatively unexplored parameter space. The launch of the GALEX satellite with its 1:25 diameter field of view, and limiting sensitivity per 1.5 ks visit of ~23 mag in the FUV ($\lambda_{\text{eff}} = 1539$ Å) and NUV ($\lambda_{\text{eff}} = 2316$ Å; Martin et al. 2005; Morrissey et al. 2007), enabled the discovery of UV variable sources in repeated observations over hundreds of square degrees for the first time.

The UV waveband is particularly sensitive to hot ($\approx 10^4$ K) thermal emission from such transient and variable phenomena as young core-collapse SNe, the inner regions of the accretion flow around accreting supermassive black holes (SMBHs), and the flaring states of variable stars. The characteristic timescales of variables and transients in the UV range from minutes to years. M dwarf flare stars have strong magnetic activity that manifests itself in flares of thermal UV emission on the timescale of minutes (Kowalski et al. 2009). RR Lyrae stars have periodic pulsations which drive temperature fluctuations from ~6000 to 8000 K that cause periodic variability in the UV on a timescale of ~0.5 d (Wheatley et al. 2012). Core-collapse SNe remain bright in the UV for hours up to several days following shock breakout, depending on the radius of the progenitor star (Nakar & Sari 2010; Rabinak & Waxman 2011), and the presence of a dense wind (Ofek et al. 2010; Chevalier & Irwin 2011; Svirski et al. 2012). Active galactic nuclei (AGNs) and quasars demonstrate stronger variability with decreasing wavelength and longer timescales of years (Vanden Berk et al. 2004).

UV variability studies of GALEX data observed as part of the All-Sky, Medium, and Deep Imaging baseline mission surveys (AIS, MIS, DIS) from 2003 to 2007, yielded the detection of M-dwarf flare stars (Welsh et al. 2007), RR Lyrae stars, AGN, and quasars (Welsh et al. 2005, 2011; Wheatley et al. 2008), and flares from the tidal disruption of stars around dormant SMBHs (Gezari et al. 2006, 2008a, 2009). Serendipitous overlap of four GALEX DIS fields with the optical Canada–France–Hawaii Telescope (CFHT) Supernova Legacy Survey, enabled the extraction of simultaneous optical light curves from image differencing for two of the tidal disruption
event (TDE) candidates (Gezari et al. 2008a), and enabled the association of transient UV emission with two Type IIP SNe within hours of shock breakout (Schawinski et al. 2008; Gezari et al. 2008b). Chance overlap of GALEX observations with the survey area of the optical Palomar Transient Factory (PTF) detected a Type IIn SN whose rising UV emission over a few days was interpreted as a delayed shock breakout through a dense circumstellar medium (Ofek et al. 2010).

Motivated by the promising results from the analysis of random repeated GALEX observations, and the demonstrated value of overlap with optical time domain surveys, we initiated a dedicated GALEX Time Domain Survey (TDS) to systematically study UV variability on timescales of days to years with multiple epochs of NUV images observed with a regular cadence of 2 days. The GALEX TDS fields were selected to overlap with the Pan-STARRS1 Medium Deep Survey (PS1 MDS; Kaiser et al. 2010). GALEX TDS and PS1 MDS are well matched in field of view, sensitivity, and cadence (shown in Table 1). Here we present the analysis of 42 GALEX TDS fields which intersect with the PS1 MDS footprint, for a total area on the sky of 39.91 deg^2, which were monitored over a baseline of 3.32 years (2008 February–2011 June). In this paper we use PS1 MDS deep stack catalogs to characterize the optical hosts of UV-variable sources, including archival optical data, a deep PS1 MDS catalog, and archival redshift and X-ray catalogs. In Section 5 we describe our sequence of steps for classifying GALEX TDS sources. In Section 7 we conclude with implications for future surveys.

### Table 1

| Survey     | Field of View (') | Plate Scale (arcsec pixel^{-1}) | PSF FWHM (') | m_{lim} per Epoch (mag) | Cadence (days) | Seasonal Visibility (months) |
|------------|-------------------|---------------------------------|--------------|-------------------------|---------------|-----------------------------|
| GALEX TDS  | 1.1               | 1.5                             | 5.3          | 23.3                    | 2             | ~1                          |
| PS1 MDS    | 3.5               | 0.258                           | 1.0          | 23.0                    | 3             | ~6                          |

### Table 2

| Name                      | R.A. (J2000) (') | Decl. (J2000) (') | E(B − V) (mag) | Epochs |
|---------------------------|-----------------|------------------|----------------|--------|
| PS_XMMLSS_MOS00           | 35.580          | −3.140           | 0.028          | 27     |
| PS_XMMLSS_MOS01           | 36.300          | −3.490           | 0.025          | 27     |
| PS_XMMLSS_MOS02           | 35.000          | −3.950           | 0.020          | 14     |
| PS_XMMLSS_MOS03           | 35.875          | −4.250           | 0.026          | 27     |
| PS_XMMLSS_MOS04           | 36.900          | −4.420           | 0.026          | 27     |
| PS_XMMLSS_MOS05           | 35.200          | −5.050           | 0.022          | 26     |
| PS_XMMLSS_MOS06           | 36.230          | −5.200           | 0.027          | 24     |
| PS_CDFS_MOS00             | 53.100          | −27.800          | 0.008          | 114    |
| PS_CDFS_MOS01             | 52.012          | −28.212          | 0.008          | 30     |
| PS_CDFS_MOS02             | 53.124          | −26.802          | 0.009          | 29     |
| PS_CDFS_MOS03             | 54.165          | −27.312          | 0.012          | 30     |
| PS_CDFS_MOS04             | 52.910          | −28.800          | 0.009          | 30     |
| PS_CDFS_MOS05             | 52.111          | −27.276          | 0.010          | 30     |
| PS_CDFS_MOS06             | 53.970          | −28.334          | 0.010          | 6      |
| PS_COSMOS_MOS21           | 150.500         | +3.100           | 0.023          | 15     |
| PS_COSMOS_MOS22           | 149.500         | +3.100           | 0.027          | 16     |
| PS_COSMOS_MOS23           | 151.000         | +2.200           | 0.024          | 24     |
| PS_COSMOS_MOS24           | 150.000         | +2.200           | 0.020          | 26     |
| PS_COSMOS_MOS25           | 149.000         | +2.300           | 0.023          | 13     |
| PS_COSMOS_MOS26           | 150.500         | +1.300           | 0.026          | 26     |
| PS_COSMOS_MOS27           | 149.500         | +1.300           | 0.019          | 27     |
| PS_GROTH_MOS01            | 215.600         | +54.270          | 0.011          | 17     |
| PS_GROTH_MOS02            | 213.780         | +54.350          | 0.015          | 16     |
| PS_GROTH_MOS03            | 214.146         | +53.417          | 0.009          | 17     |
| PS_GROTH_MOS04            | 212.400         | +53.700          | 0.011          | 18     |
| PS_GROTH_MOS05            | 215.500         | +52.770          | 0.008          | 19     |
| PS_GROTH_MOS06            | 214.300         | +52.550          | 0.008          | 8      |
| PS_GROTH_MOS07            | 212.630         | +52.750          | 0.009          | 19     |
| PS_GROTH_MOS08            | 214.300         | +52.850          | 0.009          | 8      |
| PS_GROTH_MOS10            | 242.510         | +55.980          | 0.007          | 17     |
| PS_GROTH_MOS11            | 244.570         | +55.180          | 0.009          | 17     |
| PS_GROTH_MOS12            | 242.900         | +55.000          | 0.008          | 18     |
| PS_GROTH_MOS13            | 241.300         | +55.350          | 0.007          | 18     |
| PS_GROTH_MOS14            | 243.960         | +54.200          | 0.010          | 19     |
| PS_GROTH_MOS15            | 242.400         | +54.000          | 0.011          | 21     |
| PS_GROTH_MOS16            | 241.380         | +54.450          | 0.010          | 19     |
| PS_VVDS22H_MOS00          | 333.700         | +1.250           | 0.040          | 39     |
| PS_VVDS22H_MOS01          | 332.700         | +0.700           | 0.046          | 35     |
| PS_VVDS22H_MOS02          | 334.428         | +0.670           | 0.057          | 38     |
| PS_VVDS22H_MOS03          | 333.600         | +0.200           | 0.058          | 27     |
| PS_VVDS22H_MOS04          | 334.610         | −0.040           | 0.093          | 24     |
| PS_VVDS22H_MOS05          | 333.900         | −0.720           | 0.102          | 35     |
| PS_VVDS22H_MOS06          | 332.900         | −0.400           | 0.113          | 33     |
Figure 1. GALEX TDS 1° diameter field pointings shown in blue, and the PS1 MDS 3.5° diameter field pointings shown in green. Orange hatched regions show the coverage of SDSS in the optical, red hatched regions shows the coverage of SWIRE in the optical, and cyan rectangles indicate the coverage of the CFHTLS Deep (solid lines) and Wide (dashed lines) surveys in the optical. Light gray regions show the coverage of XMM-Newton X-ray observations, and dark gray regions show the coverage of Chandra X-ray observations. (A color version of this figure is available in the online journal.)

February to 2011 June. PS1 began taking commissioning data of the MDS fields in 2009 May, but did not begin full survey operations until a year later. The GALEX far-UV (FUV) detector became non-operational in 2009 May, and so we only include near-UV (NUV) images in our study.

3. STATISTICAL MEASUREMENTS

3.1. Selection of Variable Sources

Since most galaxies are unresolved by the GALEX NUV 5″ FWHM point-spread function (PSF), we can use simple aperture photometry instead of image differencing to measure variability. We create a master list of unique source positions from the pipeline-generated catalogs (Morrissey et al. 2007) for all the individual epochs, as well as deep stacks of all the epochs, using a clustering radius of 5″. This radius is chosen such that for the typical astrometry error of GALEX of 0.5″, the Bayesian probability that the match is real is larger than the Bayesian probability that the match is spurious (Budavári & Szalay 2008). The final master list includes 419,152 sources.

In order to select intrinsically variable sources in our survey, we first need to characterize the photometric errors. Although the GALEX images are Poisson-limited, the Poisson error underestimates the total error in the GALEX catalog magnitudes by a factor of ~2 (Trammell et al. 2007). This discrepancy is attributed to systematic errors such as uncertainties in the detector background and flat-field. Thus, we measure the photometric error empirically by calculating the standard deviation of aperture magnitudes in bins of mean magnitude, \( \langle m \rangle \). We only include objects in the pipeline-generated catalogs that are detected in all or \( \geq 10 \) epochs. In each bin of \( N \) objects with \( \langle m \rangle_i = (\sum_{k=1}^{n} m_{i,k})/n \) (each bin typically has \( N = 50 \) to 1000 sources), we calculate for each epoch \( k \) of a total of \( n \) epochs,

\[
\sigma(\langle m \rangle, k) = \sqrt{\frac{1}{N - 1} \sum_{i=1}^{N} (m_{i,k} - \langle m \rangle_i)^2},
\]  

where \( m_{i,k} \) is the magnitude (in the AB system) given by \( m_{i,k} = -2.5 \log(f_6) + zp + C_{ap} - 8.2E(B-V) \), \( f_6 \) is the background-subtracted flux in a 6″ radius aperture, \( zp = 20.08 \), the aperture correction is \( C_{ap} = -0.23 \) mag (Morrissey et al. 2007), and we correct for Galactic extinction using the values for \( E(B-V) \) listed in Table 2. We use 3\( \sigma \) clipping to remove outliers in the calculation of \( \sigma(\langle m \rangle, k) \) which can arise from artifacts.

The astrometric precision depends on the signal-to-noise of the source, thus we also empirically measure a magnitude-dependent clustering radius. We do so by measuring the cumulative distribution of spatial separations between the position in each epoch and the mean position for sources in bins of \( \langle m \rangle \), and record \( d_{95}(\langle m \rangle, k) \), the value for which 95% of sources have a separation less than or equal to that value. The resulting value for \( d_{95} \) is a strong function magnitude, increasing from \( \sim 1″ \) for
Figure 2. Dates of GALEX TDS NUV observations compared to PS1 MDS observations in the \(g_{\text{P1}}, r_{\text{P1}}, i_{\text{P1}}, z_{\text{P1}}, \) and \(y_{\text{P1}}\) bands. Dotted lines show yearly intervals.

\(<m> = 18.0\) mag to \(\sim 4''\) for \(<m> = 23.0\) mag. Figures 3 and 4 show \(\sigma(<m>,k)\) and \(d_{95(<m>,k)}\) for an example GALEX TDS field PS_COSMOS_MOS23, a quadratic fit to the median function for all epochs in that field, and the median function fit over all fields.

In our master list of source positions, we include all sources detected, including sources detected in only one epoch, and fix the centroid to the epoch for which the source is detected with maximum flux. We measure forced aperture magnitudes at the positions of each source in epochs where the source was not detected by the pipeline or the spatial separation of the matched source is greater than \(d_{95(<m>,k)}\). When the aperture magnitude is fainter than \(m_{\text{lim}}\) in an epoch, it is flagged as an upper limit and replaced with \(m_{\text{lim}}\), where \(m_{\text{lim}} = -2.5 \log(5 \sqrt{(B_{\text{sky}} N_{\text{pix}} T_{\text{exp}}^2)} + zp + C_{\text{ap}}\), where \(B_{\text{sky}} = 3 \times 10^{-3}\) counts s\(^{-1}\) pixel\(^{-1}\), \(N_{\text{pix}} = 16\pi\), and \(T_{\text{exp}}\) is the exposure time of that epoch in seconds.

We select sources that have at least one epoch for which \(|m_k - <m>| > 5\sigma(<m>,k)\), where \(<m>\) is calculated only from epochs that have a magnitude above the detection limit of that epoch. We use this selection method to be sensitive to short-term and long-term variability, as well as transients. This \(5\sigma\) selection is quite conservative, and requires variability amplitudes increasing from \(|\Delta m| > 0.1\) mag for \(<m> \sim 18\) mag up to \(|\Delta m| > 1.0\) mag for \(<m> \sim 23\) mag.

Figure 3. Empirical determination of \(1\sigma\) photometric errors as a function of mean magnitude from the standard deviation of sources detected by the pipeline catalogs in \(\geq 10\) epochs for one of the GALEX TDS fields PS_COSMOS_MOS23. Solid green line shows a quadratic fit to the error function for the field PS_COSMOS_MOS23, and dashed blue line shows a quadratic fit to the median error function for all of the GALEX TDS fields. Solid red line shows the expected \(1\sigma\) Poisson error, which underestimates the total photometric error by a factor of \(\sim 2\).
We make the following cuts to the 5σ variable source sample to remove artifacts.

1. We remove sources with pipeline artifact flags indicating window bevel reflections or ghosts from the dichroic beam splitter.
2. We remove the brightest objects, with \( m \) < 18.0, due to the large area subtended by the PSF which causes uncertainty in the background subtraction.
3. We select sources within a radius of 0.55 of the center of the field, in order to avoid glints and PSF distortions, which are more prominent on the edges of the image, from un-corrected spatial distortions of photons recorded by the detectors.
4. We do not include objects that are within 1.5 of a \( m < 17 \) mag source, to avoid regions affected by the bright source’s PSF and ghost artifacts. Ghost artifacts can appear within 30′–60′ above and below a bright source in the Y detector direction. Ghosts are point-like, and thus can only be identified from their Y detector position relative to a bright source. While ghosts do not usually appear in the GALEX pipeline catalogs, we apply this cut since our forced aperture photometry could mistake ghosts for transient sources.
5. We veto objects for which in the epoch of maximum \( |m_i - m_j|/\sigma(m, k) \) or maximum flux, the aperture flux ratio of the object has \( f_6/f_{3.8} > R_{00} \), where \( f_6 \) is the 6′0 radius aperture flux, \( f_{3.8} \) is the 3′8 radius aperture flux, and \( R_{00} \) is the maximum cumulative aperture flux ratio measured for 90% of the sources in the reference source sample used to calculate \( \sigma(m, k) \) in that epoch. This cut removes fluctuations in the background due to reflections from bright stars just outside the field-of-view, as well as epochs where the PSF is distorted due to a degradation in resolution which sometimes occurs in the Y detector direction. This also vetoes cases when the pipeline shrugs a source into multiple sources, and the source is detected as variable because the center of the aperture is off-center from the peak source flux.
6. Finally, we visually inspect all of the remaining variable sources to remove any remaining artifacts that passed through the cuts above.

Figure 5 shows a gallery of good sources and vetoed variable sources from our automated cuts (for a bright reflection in panel (a), an off-center source in panel (b), and a likely ghost in panel (c)) and manual cuts (for a diffuse reflection in panel (d)). Our final GALEX TDS 5σ variable sample after the cuts listed above has a total of 1078 sources.

3.2. Variability Statistics

We characterize the variability of each 5σ variable UV source using several statistical measures. We measure the structure function (following di Clemente et al. (1996)),

\[
V(\Delta t) = \sqrt{\frac{\pi}{2} \left( \langle \Delta m_{ij} \rangle_{\Delta t} \right)^2 - \langle \sigma^2_i \rangle_{\Delta t}},
\]

where brackets denote averages for all pairs of points on the light curve of an individual source with \( i < j \) and \( t_j - t_i = \Delta t \). The two-day cadence of the observations combined with the seasonal visibility of the fields results in a distribution of time intervals between observations (shown in Figure 6) that fall into six characteristic timescale bins: \( \Delta t_{1d} = 2 \pm 0.5 \) d, \( \Delta t_{4d} = 4 \pm 0.5 \) d, \( \Delta t_{6d} = 6 \pm 0.5 \) d, \( \Delta t_{8d} = 8 \pm 0.5 \) d, \( \Delta t_{1yr} = 0.96 \pm 0.14 \) yr, and \( \Delta t_{2yr} = 1.96 \pm 0.04 \) yr. We measure the structure function in these six bins, and define \( S_k \) to be the maximum value of the structure function evaluated for \( \Delta t_{1d}, \Delta t_{4d}, \Delta t_{6d}, \Delta t_{8d}, \) and \( \Delta t_{1yr} \), and \( \Delta t_{2yr} \). We also measure the intrinsic variability as defined by Sesar et al. (2007), \( \sigma_{in} = \sqrt{\bar{\Sigma}^2 - \bar{\xi}^2} \), where \( \Sigma^2 = (1/n - 1) \sum_{k=1}^{n} (m_k - \langle m \rangle)^2 \) and
\[ \xi^2 = \left(1/n\right) \sum_{k=1}^{n} \sigma((m)_k)^2, \text{ and the maximum amplitude of variability, } \max(\{\Delta m\}). \]

3.3. UV Light Curve

In order to flag possible transient UV events that may be associated with a SN or TDE, we differentiate between stochastic variability and flaring variability. We identify flaring UV variability as sources that show a constant flux \( \geq 10 \) days before the peak of the light curve, and do not fade more than \( 2\sigma \) below the faintest pre-peak magnitude (measured \( \geq 10 \) days before the peak). This selection criteria is tailored to the NUV rise times observed in SNe (Gezari et al. 2008b, 2010; Brown et al. 2009; Milne et al. 2010; Ofek et al. 2010) and TDE candidates (Gezari et al. 2006, 2008a, 2009). We define constant pre-peak flux as a light curve with a reduced \( \chi^2 = \sum (m_k - \langle m \rangle)^2/(\sigma(m)_k)^2/(n - 1) \), where \( n \) is the number of epochs \( \geq 10 \) days before the peak. For those sources for which there are only upper limits \( \geq 10 \) days before the peak, \( \chi^2 \) is set to 1. Flaring sources with no detections before the peak are further labeled as transients. Sources with \( \chi^2 \geq 3 \), or that fade below \( 2\sigma \) of the faintest magnitude measured \( \geq 10 \) days before the peak are labeled as stochastically variable. We flag 116 flares, 145 transients, 595 stochastically variable and flaring variability. We identify flaring variables that can be associated with a SN or TDE, we differentiate between stochastic variability and flaring variability. We identify flaring variables that can be associated with a SN or TDE, we differentiate between stochastic variability and flaring variability. We identify flaring variables that can be associated with a SN or TDE, we differentiate between stochastic variability and flaring variability.

4. HOST PROPERTIES

4.1. Archival Optical Imaging Catalogs

We first characterize the host properties of the UV variable sources using archival optical \( u, g, r, i, \) and \( z \) photometry and morphology from matches to the Sloan Digital Sky Survey (SDSS) Photometric Catalog, Release 8 (Aihara et al. 2011; \( m_{\text{lim}} \sim 22 \) mag), the CFHTLS Deep Fields D1, D2, and D3 (\( m_{\text{lim}} \sim 26.5 \) mag) and Wide Fields W1, W3, and W4 (\( m_{\text{lim}} \sim 25 \) mag) merged catalogs version T0005,9 and the SWIRE ELAIS N1 and CDFS Region catalogs (\( m_{\text{lim}} \sim 24 \) mag; Surace et al. 2004). For sources with matches in multiple catalogs, we use the match from the deepest catalog. We convert the CFHTLS magnitudes to the SDSS system using the conversions in Regnault et al. (2009), and the SWIRE Vega magnitudes to the SDSS system using the transformations measured for stellar objects available at the INT WFS Web site.10 We then correct for Galactic extinction using the Schlegel et al. (1998) dust map values for \( E(B-V) \) listed in Table 2. Figure 7 shows the overlap of the GALEX TDS fields with the available archival optical catalogs.

4.2. Pan-STARRS1 Medium Deep Survey

The GALEX TDS fields overlap with the PS1 MDS fields MD01 (PS_XMMLSS), MD02 (PS_CDFS), MD04 (PS_COSMOS), MD07 (PS_GROTH), MD08 (PS_ELAISN1), and MD09 (PS_VVDS22H). The Pan-STARRS1 observations are obtained through a set of five broadband filters, \( (g_{\text{PS1}}, r_{\text{PS1}}, i_{\text{PS1}}, z_{\text{PS1}}, \text{ and } y_{\text{PS1}}) \). Further information on the passband shapes is described in Stubbs et al. (2010). The PS1 MD fields are observed with a typical cadence in a given filter of three days, with an observation in the \( g_{\text{PS1}} \) and \( r_{\text{PS1}} \) bands on night one, in the \( i_{\text{PS1}} \) band on night two, and the \( z_{\text{PS1}} \) band on night three, with \( y_{\text{PS1}} \)-band observations during each of three nights on either side of the Full Moon. Image differencing is performed on the nightly stacked images, reaching a typical \( 5\sigma \) detection limit of \( \sim 23.3 \) mag per epoch in the \( g_{\text{PS1}}, r_{\text{PS1}}, i_{\text{PS1}} \) bands and \( \sim 21.7 \) mag in the \( y_{\text{PS1}} \) band. Image difference detections from the PS1 Image Processing Pipeline (IPP; Magnier 2006) and an independent pipeline hosted by Harvard/CfA (Rest et al. 2005) are internally distributed to the PS1 Science Consortium as transient alerts for visual inspection and classification.

9 http://terapix.iap.fr/rubrique.php?id_rubrique=252
10 www.ast.cam.ac.uk/~wfcsur/technical/photom/colours
Deep stacks of the multi-epoch observations were generated to provide deep imaging with a 5σ point-source limiting magnitude of ∼24.9, 24.7, 24.7, 24.3, and 23.2 mag in the g_p1, r_p1, i_p1, z_p1, and y_p1 bands, respectively, and typical seeing (PSF FWHM) of ∼1′′.4, 1′′.3, 1′′.0, 1′′.0, and 1′′.0 in the five bands, respectively. The magnitudes are in the “natural” Pan-STARRS1 system, m = −2.5log(\text{flux})+m′, with a relative zero-point adjustment m′ made in each band for each individual epoch (Schlafly et al. 2012) before stacking to conform to the absolute flux calibration in the AB magnitude system (Tonry et al. 2012). We convert the PS1 magnitudes to the SDSS epoch (Schlafly et al. 2012) before stacking to conform to the magnitude (PSF FWHM) of ∼1′′.4, 1′′.3, 1′′.0, 1′′.0, and 1′′.0 in the five bands, respectively. The magnitudes are in the “natural” Pan-STARRS1 system, m = −2.5log(\text{flux})+m′, with a relative zero-point adjustment m′ made in each band for each individual epoch (Schlafly et al. 2012) before stacking to conform to the absolute flux calibration in the AB magnitude system (Tonry et al. 2012). We convert the PS1 magnitudes to the SDSS epoch (Schlafly et al. 2012) before stacking to conform to the absolute flux calibration in the AB magnitude system (Tonry et al. 2012), and correct for Galactic extinction using the Schlegel et al. (1998) dust map.

We calibrated these parameter cuts by comparing sources detected in both the PS1 MDS and archival optical catalogs. Figure 8 shows the PS1 star/galaxy separation criteria for 110,804 sources detected in both PS1 and SDSS catalogs in the PS_GROTH field, and for 169,461 sources detected in both PS1 and CFHT catalog in the PS_GROTH field, with i < 22 mag, the faintest magnitude for 96% of the optical hosts of the GALEX TDS sources, and the magnitude limit where all three catalogs are complete. Even though the CFHT catalogs are deeper than SDSS, they do not attempt to separate stars and galaxies for i > 21 mag, and classify all sources fainter than this magnitude as point sources. However, it is clear from both comparison plots, that the PS1 criterion of PSF_INST_MAG−PSF_AP_MAG < 0.04 mag does an even better job of separating the locus of stars from galaxies than both catalogs down to i ∼ 22 mag.

4.3. Archival Redshift and X-Ray Catalogs

We also take advantage of the many archival X-ray and spectroscopic catalogs available from the overlap of the GALEX TDS survey with legacy survey fields. In the PS_CDFS field we use X-ray catalogs from the 0.3 deg² Chandra Extended CDFS survey (Giacconi et al. 2002; Lehmer et al. 2005; Virani et al. 2006), and redshift catalogs from the VIMOS VLT Deep Survey (VVDS; Le Fèvre et al. 2004), and a compilation of redshift catalogs from GOODS and SWIRE.11 In the PS_XMMLSS field we use X-ray catalogs from the 5.5 deg² XMM-Newton Wide-Field Survey (Hasinger et al. 2007) and the 0.9 deg² Chandra COSMOS survey (Elvis et al. 2009), and redshifts from the Magellan COSMOS AGN survey (Trump et al. 2007, 2009), the VLT zCOSMOS bright catalog (Lilly et al. 2007, 2009), and the Chandra COSMOS Survey catalog (Civano et al. 2012). In the PS_GROTH field we use X-ray catalogs from the 0.67 deg² Chandra Extended Groth Strip (Nandra et al. 2005; Laird et al. 2009) and redshift catalogs from the DEEP2 Galaxy Redshift Survey (Newman et al. 2012). For PS_ELAISN1 we use the X-ray catalog from the 0.08 deg² Chandra ELAIS-N1 deep X-ray survey (Manners et al. 2003). Figure 1 shows the overlap of the GALEX TDS fields with the archival X-ray surveys. Finally, we also match the sources with the ROSAT All-Sky Bright Source and All-Sky Survey Faint Source catalogs (Voges et al. 1999, 2000).

5. CLASSIFICATION

We classify the GALEX TDS sources using a combination of optical host photometry and morphology, UV variability statistics, and matches with archival X-ray and redshift catalogs. Table 3 summarizes the sequence of steps we use to classify the sources, which we describe in detail below.

5.1. Cross-match with Optical Catalogs

We first cross-matched our 1078 GALEX TDS sources with the archival u, g, r, i, z catalogs described in Section 4.1 with a

11 http://www.eso.org/sci/activities/garching/projects/goods/MASTERCAT_v3.0.dat
Figure 9. Left: histogram of $r$ magnitudes of optical matches. Right: histogram of NUV $- r$ colors of optical matches for those sources detected in the NUV during their low-state (solid lines), and those sources with upper limits during their low-state (dashed lines).

| Step                        | $N_{\text{unclass}}$ | Archive | PS1 | Classification | Pt | Gal |
|-----------------------------|-----------------------|---------|-----|----------------|----|-----|
| Optical match               | 1078                  | 487     | 391 | 76             | 103| 21  |
| QSO color cut               | 1078                  | 326     |     |                |    | 326 |
| RRL color cut               | 753                   | 37      |     |                |    | 37  |
| Mdw color cut               | 716                   | 44      | 9   |                |    | 53  |
| Stellar locus color cut     | 663                   | 17      |     |                |    | 17  |
| $\delta_{uv}/\delta_{sd} \geq 3$ | 646                  | 19      | 15  | 34             |    |     |
| Stochastic UV var           | 616                   | 200     | 68  |                |    | 268 |
| X-ray/spec match            | 346                   | 8       | 37  |                |    |     |
|                            |                       | 302     | 21  | 358            | 37 | 53  |
|                            |                       |         |     | 22             | 305| 93  |
|                            |                       |         |     | 37             |    | 189 |

The matching radius of $3''$. This radius is recommended for matches between GALEX and ground-based optical catalogs (Budavári & Szalay 2008), and corresponds to a spurious match rate of only 1%–2% at the high Galactic latitudes of the GALEX TDS fields (Bianchi et al. 2011). However, we found that there was a population of “orphans” (no optical match within $3''$) that were detected in their NUV low-state, and had a match between $3''$–$4''$ with an optically identified quasar. Given the strong likelihood that these are real matches, we increased our matching radius to $4''$. We use the star/galaxy classifications from the catalogs to label sources as point sources (pt) or extended sources (ext). This results in 878/1078 optical matches (81%), with the majority of sources without matches in PS_CDFS, which is only partially covered by the archival catalogs.

We then match the sources that do not have archival optical matches to the PS1 MDS catalog described in Section 4.2, this increases the number of sources with optical photometry and morphology (albeit without the u band) to 1057/1078 (98%). Figure 9 shows a histogram of the $r$-band magnitude of the optical hosts, and the NUV $- r$ colors of the GALEX TDS sources in their low-state. The optical hosts have a distribution that peaks at $r \sim 21$ mag, over 3 mag brighter than the detection limit of PS1 MDS, and NUV $- r \sim 1$ mag. Sources not detected in their NUV low-state are shown as upper limits in the NUV $- r$ color histogram, and peak at NUV $- r > 2$ mag.

5.2. Orphans

We visually inspected the PS1 stack images at the locations of the 21 sources with no optical matches, and confirm that they are true orphan events. Furthermore, all of the orphans are undetected in their low-state in the NUV, with upper limits of NUV $> (22.3–23.1)$ mag. Thus the orphan hosts are likely distant stars or faint galaxies (i.e., dwarf galaxies or high-redshift galaxies) that are undetected during their low-state in the optical and NUV.

5.3. Color and Morphology Cuts

We first use the color and morphology of the optical hosts to classify the GALEX TDS $5\sigma$ UV variable sources. We define quasars as sources with optical point-source hosts with

$$u - g < 0.7$$

$$-0.1 < g - r < 1.0$$

in order to avoid the stellar locus and white dwarfs (Richards et al. 2002). Note that this color selection can be contaminated by cataclysmic variable stars (CVs), which overlap in color–color space with the quasar sample. Indeed, two of the sources classified by color as quasars are in fact spectroscopically confirmed CVs (VVDS22H_MOS05-05 and ELAISN1_MOS15-02). VVDS22H_MOS05-05 is ROTSE3 J221519.8-003257.2, a confirmed cataclysmic variable star with a dwarf-nova type spectrum. We observed ELAISN1_MOS15-02 with the APO 3.5 m telescope Dual Imaging Spectrograph (DIS) on 2011 May 3 and detected broad Balmer emission lines from a Galactic source, characteristic of a CV/dwarf nova spectrum. While these sources stood out easily because of their extreme magnitude of variability of $|\Delta u| > 4$ mag, shown in Figure 10, there
may be lower amplitude CV events still hiding in our quasar sample. However, given the low surface density of CVs relative to quasars in the sky (Szkołody et al. 2011), the expected contamination rate is consistent with the two CVs identified.

We define RR Lyrae stars as sources with optical point source hosts with

\[
\begin{align*}
0.75 < u - g < 1.45 \\
-0.25 < g - r < 0.4 \\
-0.2 < r - i < 0.2 \\
-0.3 < i - z < 0.3
\end{align*}
\] (Sesar et al. 2010). Note that the color cuts for quasars and RR Lyrae require the \( u \)-band, which is not available for sources with PS1-only matches. However, we define M dwarf stars as point sources with

\[
\begin{align*}
r - i > 0.42 \\
i - z > 0.24 \\
g < 22.2 \\
r < 22.2 \\
i < 21.3
\end{align*}
\] (West et al. 2011), which does not require \( u \)-band data. We classify stars on the main stellar locus as those with

\[
\begin{align*}
1.0 < u - g < 2.25 \\
0.4 < g - r < 1.0 \\
-0.2 < r - i < 2.0 \\
-0.3 < i - z < 1.0
\end{align*}
\] modified from Yanny et al. (2009). This color and morphology selection results in 37 RR Lyrae, 53 M-dwarf flare stars, 17 stars, and 325 quasars. Figure 11 shows the optical color–color diagram of the sources with optical point-source hosts and \( u \)-band data, and their classifications as quasars, RR Lyrae, M dwarfs, and stars.
Figure 13. Histogram of the ratio of the NUV structure function on timescales of years and days for sources classified as RR Lyrae (red) and quasars (blue). The dotted line shows the selection criteria of $S_y/S_d \geq 3$ used to classify sources with optical point source hosts as quasars.

(A color version of this figure is available in the online journal.)

Figure 14. Colors of archival optical matches to UV variables with extended optical hosts. Sources classified as AGN (by either stochastic UV variability, archival spectra, and/or an X-ray match) are color coded in green. Sources with matches with archival X-ray matches are circled in purple. Note that all galaxy sources with an archival X-ray match are defined as AGNs.

(A color version of this figure is available in the online journal.)

with $S_y/S_d \geq 3$. This is equivalent to a structure function power-law exponent cut of $\gamma > 0.2$, where $S(\Delta t) \propto \Delta t^\gamma$ (Hook et al. 1994; Vanden Berk et al. 2004; Schmidt et al. 2010). This structure-function ratio selection results in the classification of another 30 quasars. Two additional sources with optical point-source hosts have archival quasar spectra, resulting in a final quasar sample of 358. We define AGNs as sources with optically extended hosts that show stochastic UV variability (see Section 3.3), have an X-ray catalog match, and/or an archival spectroscopic classification. Figure 14 shows the optical color–color diagram of the sources with optically extended hosts, and those classified as AGN. This results in a sample of 305 AGNs. We also add archival spectroscopic classifications for six stars. This yields a total of 776/1078 (72%) sources classified as an active galaxy (quasar or AGN) or variable star.

5.5. X-Ray Sources

The archival X-ray catalogs overlap with $\sim 8.45$ deg$^2$ of the GALEX TDS survey area. Within this area, 81/89 quasars, 92/105 AGNs, and 8/9 M-dwarf stars are detected in the X-rays. In addition, there are nine optical point sources with X-ray matches that are likely quasars and M dwarfs just outside the quasar and M-dwarf color–color selection regions. UV variability selection appears to be selecting a similar population of active galaxies and M dwarfs as X-ray detection, since $\sim 90\%$ of the UV variability-selected active galaxy and M dwarf sample is also detected in the X-rays. However, only 2% of all the X-ray sources (the majority of which are active galaxies) are detected as UV variable at the selection threshold of the GALEX TDS catalog.

5.6. GUVV Catalog

We also cross-match our GALEX TDS sample with the first and second GALEX Ultraviolet Variability Catalogs (GUVV-1 and GUVV-2) from Welsh et al. (2005) and Wheatley et al. (2008). These catalogs include 894 UV variable sources ($\Delta m > 0.6$ mag in the NUV) selected from an analysis of archival GALEX AIS, MIS, DIS, and Guest Investigator (GI) fields with repeated observations. With a cross-matching radius of 4″, we find a match with 36 GUVV sources. For the 15 matches that have GUVV classifications, are all classified by GUVV as active galaxies (AGN or quasar), which are in agreement with our GALEX TDS classifications. Of the 21 matches without GUVV classifications, we find 5 sources classified by GALEX TDS as RR Lyrae, 9 classified as active galaxies (AGN or quasar), 6 with optical point-source hosts, and 1 with a galaxy host.

5.7. Unclassified Sources

The remaining 302 unclassified sources include 91 with optical point-source hosts, which are likely stars, quasars with non-standard colors (high-redshift or reddened), or unresolved galaxies, 190 with galaxy hosts, and 21 orphans. The 190 galaxy hosts may either be faint AGNs with poorly constrained UV light curves, or hosts of UV-bright extragalactic transients. In Figure 15 we show the maximum $|\Delta m|$ in the NUV as a function of low-state NUV magnitude of the remaining unclassified sources. The unclassified UV source with the most extreme amplitude, ELAISN1_MOS15-09 with $|\Delta m| > 4.2$ mag, was...
spectroscopically confirmed to be from the nucleus of an inactive galaxy at $z = 0.1696$, and its UV/optical flare detected by GALEX TDS and PS1 MDS (PS1-10jh) was attributed to the tidal disruption of a star around an SMBH (Gezari et al. 2012). Also in this sample is a UV transient spectroscopically confirmed to be a Type IIP SN 2010aq at $z = 0.086$ (COSMOS_MOS26-29), whose UV/optical light curve from GALEX TDS and PS1 MDS was fitted with early emission following SN shock breakout in a red supergiant star (Gezari et al. 2010). Both of these spectroscopically classified extragalactic transients are labeled in Figure 15.

Our 5σ selection criterion translates to a limiting sensitivity to transients in a host galaxy with a magnitude $m_{\text{host}}$ of a magnitude of

$$m_{\text{trans}} = -2.5 \log \left( \frac{10^{-5.5} - 10^{-m_{\text{host}}}}{5} \right),$$

which ranges from $m_{\text{trans}} \sim 20.0$ mag for $m_{\text{host}} = 18$ mag to $m_{\text{trans}} \sim 22.7$ mag for $m_{\text{host}} = 23$ mag. Thus, our variability selection threshold is less sensitive to transients in host galaxies with bright NUV fluxes. On the red sequence of galaxies, where $M_{\text{NUV}} \approx -14.5$ (Wyder et al. 2007), this selection effect is not as much of an issue, since already for $z > 0.05$ one gets $m_{\text{host}} > 22$ mag. However, star-forming galaxies on the blue sequence are 2.5 mag brighter in the NUV, and thus the host galaxy brightness can be a factor in reducing the sensitivity to faint transients. For example, our GALEX TDS 5σ sample does not include SN 2009kf, a luminous Type IIP SN in a star-forming galaxy at $z = 0.182$ which we reported our GALEX TDS detection of in Botticella et al. (2010). This source varied at only the 4.25σ level in the NUV during its peak. However, because this transient was selected from a spatial and temporal coincidence with a PS1 transient alert, we could lower our threshold for variability selection in the UV. The systematic selection of SN and TDE candidates from the joint GALEX TDS and PS1 MDS transient detections will be presented in future papers.

6. DISCUSSION

6.1. Classification Demographics

Figure 16 shows a pie diagram of the source classifications. Out of the total of 1078 GALEX TDS sources, 62% are classified as actively accreting SMBHs (quasars or AGN), and 10% as variable and flaring stars (including RR Lyrae, M dwarfs, and CVs). Note that the relative fraction of the different classes of sources is sensitive to both their intrinsic magnitude distribution, and the magnitude-dependent variability selection function of the sample. Table 4 gives the GALEX TDS catalog, sorted by decreasing NUV amplitude, with the GALEX ID, R.A., decl., low-state NUV magnitude, maximum amplitude of NUV variability ($\Delta|m_{\text{max}}|$), intrinsic variability ($\sigma_{\text{int}}$), the structure function on day ($S_d$) and year ($S_y$) timescales; the characteristics of the NUV light curve: flaring (F) or stochastically variable (V); the morphology of the matching optical host: point-source (pt) or extended (ext); the color classification of the matching optical host: RR Lyrae (RRL), M dwarf star (Mdw), star (Star), or quasar (QSO); the archival redshift, an X mark if there is a match with an archival X-ray source; and finally the GALEX TDS classification: RR Lyrae (RRL), M dwarf star (Mdw), star (Star), quasar (QSO), AGN, UV flaring source or UV transient source with galaxy host (Galaxy), no host (Orphan), and point-source hosts (Point).

(A color version of this figure is available in the online journal.)
AGNs, and RR Lyrae, respectively. For the transient sources, we can calculate a total surface density rate, #/area × t_{eff}, where t_{eff} is the effective survey time at the cadence that matches the characteristic timescale of the transient. For extragalactic transients such as young SNe and TDEs, which vary on a timescale of days, we use the time intervals for which the fields were observed with a cadence of 2.0 ± 0.5 days. If we include all flaring or transient GALEX TDS sources with a galaxy host, this yields a surface density rate of 52 ± 38 deg^{-2} yr^{-1} for extragalactic transients. For M dwarfs which vary on timescales shorter than an individual observation, we use the total exposure time for each epoch. If we assume a survey with a cadence of two days and t_{exp} = 1.5 ks, this translates to a surface density rate for M dwarfs of 15 ± 10 deg^{-2} yr^{-1}.

6.2. UV Variability Properties of Classified Sources

Various optical studies of rest-frame UV variability in high redshift quasars have demonstrated that the variability of AGN increases with decreasing rest wavelength (di Clemente et al. 1996; Vanden Berk et al. 2004; Wilhite et al. 2005). In Figure 18, we show histograms of $\sigma_{int}$ for the UV variable sources with classifications. Quasars show a distribution of $\sigma_{int}$ with a mean that is ∼2 times larger than measured at optical wavelengths from the SDSS Stripe 82 sample from Sesar et al. (2007). This effect is even more pronounced in the magnitude of the structure function on years timescales ($\sigma_{z}$), which has a mean that is five times larger than the mean measured in the r-band ($\lambda_{eff} = 6231$) from Schmidt et al. (2010). This trend is consistent with the wavelength-dependent rise in variability amplitude observed in the structure function for quasars in the rest-frame UV (Vanden Berk et al. 2004) and observed UV (Welsh et al. 2011). The fact that AGN become bluer during high states of flux (Giveon et al. 1999; Geha et al. 2003; Gezari et al. 2008a) has been attributed to increases in the characteristic temperature of the accretion disk in response to increases in the mass accretion rate ($\dot{M}$) in accretion disk models. In a future study, we will use simultaneous UV and optical light curves from GALEX TDS and PS1 MDS for our 358 individual quasars to test this result with a larger dynamic range in wavelength.

For the subsample of 95 quasars with archival redshifts ($z_{\text{mean}} = 1.26$, $\sigma_{z} = 0.39$), in Figure 19 we plot $\sigma_{int}$ versus the low-state NUV absolute magnitude, and find a steep negative correlation fitted by $\log(\sigma_{int}) = (1.6 \pm 0.1) + (\beta/2.5)M_{\text{NUV}}$, where $\beta = 0.24 \pm 0.04$, in excellent agreement with the trend for increased variability in lower luminosity quasars seen from optical observations with $\beta = 0.246 \pm 0.005$ (Vanden Berk et al. 2004), and shallower than expected for a Poissonian process which has $\beta = 0.5$ (Cid Fernandes et al. 2000). We also show the subset of 68 AGN with archival redshifts ($z_{\text{mean}} = 0.64$, $\sigma_{z} = 0.55$), which clearly do not show a relation between $\sigma_{int}$ and low-state NUV absolute magnitude. This is most likely a result of dilution of the variability amplitude from the contribution of the host galaxy in the NUV.

The largest values of $|\Delta m|$ (plotted in Figure 10) are found in RR Lyrae and M dwarfs, with a tail of large amplitude variations reaching up to $|\Delta m| = 2.9$ mag in RR Lyrae and up to $|\Delta m| = 4.6$ mag in M dwarfs. In the optical, the RR Lyrae structure function is very weakly dependent on timescale, with an amplitude of 0.1–0.2 mag (Schmidt et al. 2010). The NUV structure function also shows a weak dependence of amplitude on timescale when comparing the structure function on days to years timescales, however, with an amplitude that is
photometric errors of $\sigma(m) \sim 0.01$ mag. In coordination with a ground-based optical survey, such as Pan-STARRS2 (Burgett 2012) or LSST,\textsuperscript{12} this could yield the simultaneous UV and optical detection of $\approx 10^4$ variable quasars and $\approx 10^8$ RR Lyrae and M dwarfs, as well as increase the discovery rate of UV-bright extragalactic transients (young SNe and TDEs) by a factor of $\approx 100$.

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