Linear polarization of rapidly rotating ultracool dwarfs

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Abstract. We present imaging linear polarimetry data of a sample of 18 late-M, L-, and early-T type dwarfs taken with the $Z$- and $J$-band filters and the LIRIS instrument of the 4.2-m William Herschel Telescope. All of our targets have projected rotational velocities $\geq 30$ km s$^{-1}$ and oblate ultracool atmospheres ($T_{\text{eff}} < 2700$ K), which may harbor clouds of condensate particles. Our polarimetric measurements have typical error bars of $\pm 0.13\%$, i.e., linear polarization degrees larger than 0.4% can be detected with a confidence of $\geq 3\sigma$. Seven dwarfs appear to be polarized in the $J$-band with indices of $P = 0.4$–$0.7\%$, suggesting the presence of atmospheric dusty structures. There is a hint that the dwarfs with the largest rotations ($v \sin i \geq 60$ km s$^{-1}$) show higher incidence of detected $J$-band linear polarization than the dwarfs with smaller projected rotational velocities. We also detect linear polarization variability indicative of “weather”.

Key words. polarization – brown dwarfs – stars: atmospheres – stars: late-type – stars: low-mass

1. Introduction

Dwarfs with spectral types cooler than M7 ($T_{\text{eff}} \leq 2700$ K; typically referred to as ultracool dwarfs) are believed to undergo the formation of a wide range of atmospheric condensate species (solid and liquid particles) such as corundum ($\text{Al}_2\text{O}_3$), iron (Fe), enstatite ($\text{MgSiO}_3$), forsterite ($\text{Mg}_2\text{SiO}_4$), titanium dioxide ($\text{TiO}_2$), and gehlenite ($\text{Ca}_2\text{Al}_2\text{SiO}_7$) among others (Jones & Tsuji 1997; Ackerman & Marley 2001; Helling et al. 2008; Witten et al. 2011). According to models, these condensates may build up into structures (e.g., “clouds”), which are located in the outer layers of the atmosphere for $T_{\text{eff}} \geq 1300$ K and in layers below the visible photosphere for $T_{\text{eff}} \leq 1300$ K (Allard et al. 2001). Condensates or “dusty particles” represent one relevant and yet poorly understood source of opacity in ultracool dwarfs. Atmospheric dust may also polarize the object’s output light at particular wavelengths through scattering processes as suggested by the theoretical work of Sengupta & Krishan (2001), and observationally demonstrated by the detection of linear polarization at optical and near-infrared frequencies (Ménard et al. 2002; Zapatero Osorio et al. 2003).
2005; Goldman et al. 2009; Tata et al. 2009; Zapatero Osorio et al. 2011). Polarization may become a useful tool to comprehend the complexity of ultracool atmospheres.

Additionally, ultracool dwarfs have high values of projected rotational velocities \(v \sin i\) indicating that they are indeed rapid rotators (Zapatero Osorio et al. 2006; Konopacky et al. 2012, and references therein).

This fact combined with convection and the presence of dust could give rise to intricate atmospheric dynamics, likely generating periodic and non-periodic photometric variability as seen in some late-M, L, and T dwarfs (Bailer-Jones & Mundt 2001; Martin et al. 2001; Koen 2003; Buenzli et al. 2012).

From a theoretical perspective, Sengupta & Krishan (2001), Sengupta (2003), Sengupta & Kwok (2008), Sengupta & Marley (2010), and de Kok et al. (2011) predicted that ultracool dwarfs with atmospheric condensates and high \(v \sin i\)’s show measurable linear polarization degrees of typically \(\leq 1\%\) in the optical and near-infrared. Fast rotation makes dwarf to take the shape of an oblate ellipsoid, and this lack of symmetry leads to incomplete cancellation of the polarization from different areas of the surface. Gravity is an additional ingredient to take into account since rotationally induced non-sphericity is favored at lower atmospheric gravities. For a similar amount of dust particles in the ultracool atmospheres it is expected that the largest rotations and lowest gravities produce the largest polarization degrees.

Here, we aim at studying the capabilities of linear polarization in the near infrared to probe the presence of atmospheric condensates in ultracool dwarfs and to shed new light on the dependence of the linear polarization with rotation. The results shown in this proceeding have been recently accepted for publication (Miles-Páez et al. 2013).

2. Target selection and observations

We selected 18 bright \((10.3 < J < 14.9\) mag) ultracool dwarfs with spectral types from M7 through T2 and published projected rotational velocities \(v \sin i \geq 30\) km s\(^{-1}\) (Zapatero Osorio et al. 2006; Konopacky et al. 2012, and references therein). All dwarfs are observable from northern astronomical observatories and sufficiently bright at near-infrared wavelengths to achieve accurate polarimetric photometry \((\sigma \leq 0.2\%)\) using short exposures and 4-m class telescopes. They represent \(\sim 37\%\) of all dwarfs cooler than M7 that have \(v \sin i \geq 30\) km s\(^{-1}\) available in the literature (see Figure 1).

In Table 1, we provide the targets complete names and their spectral types.

We carried out several linear polarimetric campaigns using the Z and J-bands filter and LIRIS on the 4.2–m William Herschel Telescope in La Palma island (Spain). In Table 1, we provide the filters and dates of observations for each target. These were observed following a nine–point dither pattern for a proper sky background contribution removal. Typical individual exposure times were in the range 5–180 s depending on the target’s brightness, filter and seeing.

In its polarimetric imaging mode, LIRIS uses a Wedged double Wollaston device, consisting in a combination of two Wollaston prisms that deliver four simultaneous images of the polarized flux at vector angles 0°, 90°, 45° and 135°. Raw data frames were divided into four slices corresponding to the po-
Table 1. Targets and linear polarization measurements

| Name                  | SpT | Fil | Obs. date (JD-2450000.5) | $p$ (%) | $\theta$ (deg) |
|-----------------------|-----|-----|--------------------------|---------|---------------|
| 2MASS J00192626+4614078 | M8  | J   | 6206.8719                | 0.38 ± 0.15 | –             |
|                       |     | Z   | 6207.9954                | 0.57 ± 0.13 | 146.8 ± 6.7   |
| BRI 0021–0214         | M9.5| J   | 6206.8988                | 0.09 ± 0.11 |              |
| LP 349–25AB           | M8+M9| J   | 6206.9830                | 0.18 ± 0.11 |              |
|                       |     | Z   | 6208.4696                | 0.20 ± 0.10 | –             |
| 2MASS J00361617+1821104 | L3.5| J   | 6207.0147                | 0.00 ± 0.11 | –             |
| 2MASS J00452143+1634446 | L2  | J   | 6207.0275                | 0.00 ± 0.11 | –             |
| 2MASS J02281101+2537380 | L0  | J   | 6207.0697                | 0.35 ± 0.14 | –             |
| LP 415–20AB           | M7+M9.5| J   | 6207.2099                | 0.39 ± 0.12 | 174.9 ± 9.0   |
|                       |     | Z   | 6208.2221                | 0.38 ± 0.12 | 176.3 ± 9.4   |
| 2MASS J07003664+3157266AB | L3.5+L6| J   | 6207.2425                | 0.48 ± 0.14 | 87.3 ± 7.8    |
| 2MASS J08283419–1309198 | L2  | J   | 5927.1754                | 0.14 ± 0.10 | –             |
|                       |     | J   | 6321.0538                | 0.22 ± 0.15 |              |
| 2MASS J11593850+0057268 | L0  | J   | 6094.9304                | 0.53 ± 0.15 | 28.3 ± 8.0    |
|                       |     | J   | 6321.1119                | 0.40 ± 0.13 | 179.0 ± 9.0   |
| 2MASS J12545393–0122474 | T2  | J   | 6321.1575                | 0.00 ± 0.34 | –             |
| 2MASS J14112131–2119503 | M9  | J   | 6094.9835                | 0.48 ± 0.17 | 148.7 ± 10.0  |
| 2MASS J15010818+2250020 | M0  | J   | 6321.1950                | 0.51 ± 0.12 | 46.2 ± 7.0    |
| 2MASS J15210103+5053230 | M7.5| J   | 6321.2133                | 0.60 ± 0.13 | 13.8 ± 6.2    |
| 2MASS J18071593+5015316 | L1.5| J   | 6094.1371                | 0.67 ± 0.11 | 67.0 ± 5.0    |
|                       |     | Z   | 6207.8720                | 0.15 ± 0.15 | –             |
| 2MASS J18353790+3259545 | M8.5| J   | 6094.1755                | 0.07 ± 0.10 | –             |
| 2MASS J20360316+1051295 | L3  | J   | 6094.1880                | 0.18 ± 0.16 | –             |
|                       |     | Z   | 6207.9429                | 0.73 ± 0.28 | –             |
| 2MASS J20575409–0252302 | L1.5| J   | 6206.8253                | 0.40 ± 0.15 | –             |

The linear polarization vectors, and each slice was reduced following standard procedures for the near infrared (sky subtraction, flat-field correction, alignment and combination of the nine images of the pattern).

3. Results and discussion

Linear polarization degrees and polarization angles were derived from the Stokes parameters as explained in Zapatero Osorio et al. (2011). The $J$ magnitudes of the fast rotators sample are in the range 10.3-14.1 mag, which provide enough photons to reach accuracies of ±0.13% in the $Z$ and $J$ linear polarimetric measurements. According to the $3\sigma$ criterion, $P/\sigma \geq 3$ ($P$ and $\sigma$ are linear polarization degree and its associated uncertainty), we can detect polarimetric signals with indices higher than $P = 0.4\%$. About 40% of our sample (7 out of 18 sources in the $J$-band, and 2 out of 5 in the $Z$ filter) appear to be linearly polarized.

Because the linear polarization degree $P$ is always a positive quantity, small values of $P$ and values of $P$ affected by poor signal-to-noise-ratio data are statistically biased toward an overestimation of the true polarization. We applied the equation given by Wardle & Kronberg (1974) to derive the debiased linear polarization degree, $p^*$, by taking into account the measured $P$ and its associated uncertainty:

$$p^* = \sqrt{P^2 - \sigma^2}$$

(1)

The debiased linear polarization degree ($p^*$) and the polarization vibration angle ($\theta$) are provided in Table 1.

3.1. Linear polarization versus rotation

In Figure 2, we plotted the degree of linear polarization in the $J$ (left) and $Z$ (right) bands as
Fig. 2. Z-band (right) and J-band (left) degree of linear polarization as a function of $v \sin i$ for our sample of fast rotators. Polarized objects (i.e., $P/\sigma \geq 3$) are plotted as encircled symbols. Diamonds are objects taken from [Zapatero Osorio et al. (2011)], the arrow shows the upper limit for $v \sin i$.

3.2. Time variability of the linear polarization

We investigated the variability of the linear polarization degree for objects with measurements available on different occasions. We find that both the degree of linear polarization ($P$) and the vibration angle of the polarization ($\Theta$) remain almost constant (within the error bars) in consecutive measurements separated by a day or ~ 20 rotations (assuming $v_{\text{rot}} = 60$ km s$^{-1}$ and size 1 $R_{\text{Jup}}$). However, when comparing measurements of the same object taken in the same filter, using the same instrumental configuration but separated by several months (or a few thousands rotation cycles), $P$ and $\Theta$ varies. We attribute this to changes in the dusty cloud patterns of the atmospheres likely due to rapid rotation and convective motions [Showman & Kaspi (2013); de Kok et al. (2011)].

4. Conclusions

About 40±15% of rapidly rotating, field M7-L3.5 dwarfs show significant linear polarization at 1.25 μm with degrees in the interval $p^* \sim 0.4$–0.8%, in agreement with published theoretical models.
Fig. 3. J-band degree of linear polarization as a function of the rotation period. Polarized objects are plotted as encircled symbols. Diamonds stand for objects taken from Zapatero Osorio et al. (2011).

Also, we find that linear polarimetric measurements are stable over periods of a few days or a few tens rotation cycles. However, they change dramatically over periods of a few months or thousands rotation cycles. We attribute these variations to changes in the clouds due to rotation and convection.

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