ELEMENTAL ABUNDANCES IN THE X-RAY GAS OF EARLY-TYPE GALAXIES WITH XMM-NEWTON AND CHANDRA OBSERVATIONS

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ABSTRACT

The source of hot gas in elliptical galaxies is thought to be due to stellar mass loss, with contributions from supernova (SN) events and possibly from infall from a surrounding environment. This picture predicts supersolar values for the metallicity of the gas toward the inner part of the galaxy, which can be tested by measuring the gas phase abundances. We use high-quality data for 10 nearby early-type galaxy from XMM-Newton, featuring both the European Photon Imaging Camera and the Reflection Grating Spectrometer, where the strongest emission lines are detected with little blending; some Chandra data are also used. We find excellent consistency in the elemental abundances between the different XMM-Newton instruments and good consistency with Chandra. Differences in abundances with aperture size and model complexity are examined, but large differences rarely occur. For a two-temperature thermal model plus a point source contribution, the median Fe and O abundances are 0.86 and 0.44 of the solar value, while Si and Mg abundances are similar to that for Fe. This is similar to stellar abundances for these galaxies but SNe were expected to enhance the gas phase abundances considerably, which is not observed.

Key words: cooling flows – galaxies: abundances – galaxies: elliptical and lenticular, cD – galaxies: individual (NGC 720, NGC 1399, NGC 3923, NGC 4406, NGC 4472, NGC 4553, NGC 4636, NGC 4649, NGC 5044, IC 1459) – X-rays: galaxies

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1. INTRODUCTION

Early-type galaxies possess an interstellar medium that is dominated by hot gas \(3\times10^6\) K), although the mass of gas can vary widely between systems. The origin of the hot gas is not entirely a settled issue, but it is probably the result of mass loss from stars within the galaxy as well as infall onto the galaxy, especially when it lies in a galaxy group. The abundances of this gas reflect its history and can potentially inform us as to the number of Type I and Type II supernovae (SNe) that must have been present. There have been a number of surprises in the abundance measurements, such as that the abundance is lower than initially predicted for gas shed from stars and enhanced by SNe (e.g., Arimoto et al. 1997). Also, the values for the gas abundances have varied considerably, for different models applied to the same galaxy, and between galaxies, so a uniform picture has been slow to emerge.

There are a variety of issues that face investigators when determining abundances within early-type galaxies (or other systems with thermal gas). There is the problem of instrumental calibration, which can be notoriously difficult, despite dedicated efforts by the scientific staff. It is often several years after launch of a mission before most of the important calibration issues are understood. As the calibrations for XMM-Newton and Chandra have matured, this is a good time to examine the spectra of similar objects and compare the results.

Another issue is that the derived metallicities are sensitive to the number of spectral components used in a model. This was pointed out by Trinchieri et al. (1994), among others, who showed that when a single-temperature thermal model was applied to the luminous emission from an elliptical galaxy, the derived metallicity was significantly lower than when a two-temperature (2T) model was used. Often, the \(\chi^2\) is acceptable for both models, so without further information, it is difficult to identify the correct model. The problem is one of resolution, both spatially and spectrally. On the one hand, the temperature can vary within a galaxy, so by analyzing projected emission, there are multiple temperature components. Also, there are point sources, mainly due to low-mass X-ray binaries, and these provide a hard continuum that must be accounted for. Even with Chandra, not all of the individual point sources can be excluded, although their collective spectra are fairly constant from galaxy to galaxy (Irwin et al. 2003), so modeling of this component is tractable. Spectrally, one could identify the need for various spectral components if it were possible to measure lines of various ionization states for the same element. ROSAT did not have sufficient spectral resolution nor bandpass coverage to constrain many of the important parameters, nor did it have the spatial resolution to remove point sources. The ASCA satellite could measure the high energy contribution from the X-ray binaries, but had very poor spatial resolution and there were calibration issues at the important low-energy part of the detector.

Some of these issues are resolved by using Chandra and XMM-Newton observations. The Chandra data have excellent spatial resolution, so most point sources can be excluded, and the combination of spectral resolution and calibration is superior to its predecessors. In comparison, XMM-Newton has poorer spatial resolution but more collecting area and a relatively high dispersion grating spectrum is obtained for all on-axis targets.

There is a range of results that seems puzzling, as both high and low metallicities are found in optically similar galaxies. Individual XMM-Newton observations show a similar range of behavior, with subsolar abundances (referred to Anders & Grevesse 1989), as in NGC 6251 (Sambruna et al. 2004), NGC 3585, 4494, and 5322 (O’Sullivan & Ponman 2004), near-solar values, as in NGC 4649 (Randall et al. 2006), through the supersolar values seen in NGC 507 (Kim & Fabbiano 2004). Similarly, individual Chandra observations of galaxies can be of near solar metallicity, such as in NGC 1316 (Kim & Fabbiano 2003) or NGC 4649 (Randall et al. 2004), but other galaxies can show quite low abundances, such as NGC 1291 (Irwin et al.
2. Sample Selection and Data Reduction

2.1. Sample Selection

The sample is based on the work of Brown & Bregman (1998), who obtained a complete sample of 34 early-type galaxies which were optically selected and flux-limited. We selected 10 sources in their sample which were observed by XMM-Newton, and are bright enough to use both of the data of the European Photon Imaging Camera (EPIC) and the Reflection Grating Spectrometer (RGS). This is mainly determined by the quality of the RGS spectra, which were included in our sample by presenting enough emission line features in the preview spectra. Several of these galaxies also have high-quality Chandra data, which were used to determine abundances as well. The observations that were used in this paper are given in Table 1.

2.2. Chandra ACIS Data Processing

Chandra observations of NGC 4649, NGC 4472, and NGC 1399 were obtained from the HEASARC online archive (http://heasarc.gsfc.nasa.gov/docs/archive.html). The data for each observation were processed in a uniform manner following the Chandra data reduction threads using CIAO version 3.4. Times of high background were removed from the data, and all images created were corrected for exposure and vignetting. Point sources were identified and subsequently removed using “Mexican-Hat” wavelet detection routine WAVDETECT in CIAO on the 0.3–6.0 keV band image. For each galaxy, we extracted spectra from a region 1′ in diameter centered on the nucleus of the galaxy (the same region as was used with XMM-Newton data for these galaxies). Background spectra were extracted from a region near the edge of the S3 chip for each galaxy. While the chosen background region may contain extended emission from the galaxy (which typically fills the entire S3 chip), it only represents ~1% of the total emission in the 1′ diameter aperture, so precise treatment of the background is unnecessary. The spectra were grouped such that each spectral bin contained at least 25 counts.

2.3. XMM-Newton EPIC Data Reduction

The EPIC and RGS data of the 10 sources were processed with standard procedures using the XMM-Newton Science Analysis System (SAS, ver. 7.0.0).

Table 1

Properties and Observational Information of Surveyed Galaxies

| Galaxy   | Type$^a$ | $B_T^{0.8}$ (mag) | D$^b$ (Mpc) | $N_{H}^{c}$ (10$^{20}$ cm$^{-2}$) | $r_{1.7}^{d}$ (arcsec) | ObsID$^e$ | MOS1 | MOS2 | PN | RGS1 | RGS2 | ACIS-S | Net Exposure Time (ks)$^f$ |
|----------|---------|------------------|------------|-------------------------------|------------------------|----------|------|------|-----|------|------|-------|-------------------|
| NGC 720  | E       | 11.13            | 27.67      | 1.55                          | 39.87                  | 0112300101| 29.6 | 30.0 | 19.9 | 46.3 | 45.3 | 55.9 | 2455.3           |
| NGC 1399 | E0pec   | 10.44            | 19.95      | 1.31                          | 42.55                  | 0400620101| 121.5| 120.5| 73.9 | 118.2| 118.1| 115.1| 746.8            |
| NGC 1399 |         |                  |            |                               |                        | 319$^g$  |      |      |      |      |      | 32.5 |                 |
| NGC 3923 | E4-5    | 10.62            | 22.91      | 6.29                          | 53.35                  | 0027340101| 38.8 | 38.7 | 29.8 | 43.9 | 42.6 | 88.6 | 325.4            |
| NGC 4406 | S0(3)/E3| 9.74             | 17.14      | 2.58                          | 89.64                  | 0108260201| 77.9 | 78.9 | 47.8 | 83.5 | 81.1 | 183.5| 165.2            |
| NGC 4472 | E2/S0   | 10.14            | 26.29      | 1.65                          | 104.40                 | 0200130101| 82.5 | 82.7 | 72.8 | 101.5| 101.5| 71.5 | 466.2            |
| NGC 4472 |         |                  |            |                               |                        | 319$^g$  |      |      |      |      |      | 32.5 |                 |
| NGC 4552 | E        | 10.57            | 15.35      | 2.56                          | 48.89                  | 0114570101| 27.8 | 31.2 | 18.5 | 42.9 | 42.8 | 32.5 | 325.4            |
| NGC 4636 | E/S0     | 10.43            | 14.66      | 1.83                          | 100.08                 | 011190701| 59.2 | 59.3 | 51.1 | 62.9 | 61.4 | 183.5| 165.2            |
| NGC 4649 | E2      | 9.70             | 16.83      | 2.13                          | 73.73                  | 0021540201| 50.6 | 50.6 | 42.2 | 53.1 | 51.6 | 82.5 | 71.5             |
| NGC 4649 |         |                  |            |                               |                        | 78$^b$   |      |      |      |      |      | 22.9 |                 |
| NGC 5044 | E0      | 11.67            | 31.19      | 5.03                          | 82.23                  | 0037950101| 22.6 | 22.7 | 17.0 | 23.6 | 22.8 | 32.5 | 822.3            |
| NGC 5044 |         |                  |            |                               |                        | 78$^b$   |      |      |      |      |      | 22.9 |                 |
| IC 1459  | E3      | 10.83            | 29.24      | 1.19                          | 38.61                  | 0135980201| 29.3 | 29.3 | 25.2 | 31.7 | 30.8 | 22.9 | 822.3            |

Notes.

$^a$ The galaxy type was taken from NED.

$^b$ Total B-band magnitude from RC3 (De Vaucouleurs et al. 1991).

$^c$ Distances in Mpc, measured by a surface brightness fluctuation method (Tonry et al. 2001).

$^d$ The Galactic H I column density, taken from the dust map by Dickey & Lockman (1990).

$^e$ Effective, blue-half light radius in arcseconds derived from RC3 (De Vaucouleurs et al. 1991).

$^f$ XMM-Newton observation ID.

$^g$ Chandra ObsID.
For the EPIC data, high background periods (counts rate >0.35 count s⁻¹ for MOS CCDs and >1 count s⁻¹ for PN CCD) were removed by examining light curves of events with photon energy greater than 10 keV and PATTERN = 0, which is very sensitive to the X-ray flares. Photon events with PATTERN ≤ 12 were retained for the MOS cameras while PATTERN ≤ 4 were retained for the PN camera. We chose the photon events in the energy range 0.3–7.0 keV for the MOS and 0.4–7.0 keV for the PN to eliminate the calibration uncertainty at the lower energy band (<0.3 keV for MOS and <0.4 keV for PN) and the uncertainty at the high energy band (>7.0 keV) due to very low photon counts.

Two different source regions were extracted from the cleaned event lists. The regions are 1′ and 4′ in diameter, centered at the peak of the source X-ray emission, which is also the optical center of the galaxy. The reason for this is that the smaller region will be more uniform than the larger region, when radial gradients exist in the gas, such as temperature gradient. The larger region is representative of the average properties of the entire galaxy, so this analysis might be useful for comparisons with more distant galaxies. The larger diameter region will sample a broader variation of gas properties, such as in temperature and abundance, so by comparing the spectral fitting results from the two regions, we can assess the implied variations.

We subtract the background regions on the same observation fields, which are a few arcminutes away from the central sources and contain little hot gas emission from the central sources; point sources were excluded. The background regions are larger than the source regions to compensate the vignetting effect on the edge, and are circular regions with diameters of 1.67 and 5.83 for source regions of 1′ and 4′, respectively. Redistribution matrix files (RMF), which contain the instrumental responses, and ancillary response files (ARF), which contain the effective area of the detector, were generated with the SAS. The spectra were grouped so that each energy bin contains a minimum of 25 counts, appropriate for χ² fitting.

2.4. XMM-Newton RGS Data Reduction

For the RGS data, high background periods were screened by examining the light curve of CCD9 (excluding the source region), which is the most sensitive CCD to be affected by the background flares. The screening criteria are the same as employed by Tamura et al. (2003), because we are using the same template RGS background. Events were selected in the energy range of 0.33–2.48 keV (5–38 Å), and the first-order spectra of RGS1 and RGS2 were extracted.

The RGS is a slitless spectrograph with a field of view of 5′ in the cross-dispersion direction and 1′ in the dispersion direction (covering the whole diameter of the MOS field of view). The RGS1 and RGS2 are perpendicular to each other and therefore sample different projected spatial regions of the galaxies. The size of the extended emission can be inferred from the their EPIC images, and we can fit the surface brightness distribution with a “beta” model, $S_r \propto \left(1 + \frac{r}{r_c}\right)^{-\beta + 1/2}$, where $\beta \sim 0.5$ and $r_c \sim 0.2$. Although some of these galaxies are detectable in X-rays to 5′ the flux is concentrated. Based on the surface brightness distribution, we defined two source extraction regions in the cross-dispersion direction. For a cross-dispersion width of 1′ (to be compared with 1′ diameter regions), 90% of the flux is extracted, while in our larger extraction width of 3.5, 99% of the flux is extracted. These extraction regions are somewhat different from the circular apertures used to analyze the EPIC data, which are a consideration in the spectral fitting process. If there is a gradient in spectral properties, the narrower extraction width will suffer less spectral blending. The advantage of the spectrum with the larger extraction width is higher S/N, provided that the source remains brighter than the background.

When extracting a spectrum, a background must be removed and the standard procedure is to use a local background, defined as a region in the same CCD, after excluding sources. However, some of these galaxies fill the field of view of the RGS CCDs, so there is no emission-free background region that can be used. Therefore, we used a background from other observations. For our observations, we used the RGS template background (Tamura et al. 2003), combined with several other blank sky observations. The background is selected with the same region on RGS CCDs of the template as the source region. Another advantage of this method is that the background is better defined than a local background due to the very long exposure times of the background fields.

Following screening and background subtraction, the RGS spectra were binned so that each energy bin contain a minimum of 25 counts, making the data appropriate for χ² fitting.

3. SPECTRAL FITTING TECHNIQUES AND RESULTS

3.1. Spectral Fitting Techniques

Originally, we anticipated using the RGS spectral fits to yield unambiguous abundance results, as individual lines are resolved from O viii (19 Å) and highly ionized Fe, such as Fe xvii (15, 17 Å) and Fe xviii (16 Å) (see Figure 1), whereas they are badly blended in the EPIC data. These RGS data provide clear identification of the line strengths and the absence of certain important lines, such as the O vii triplet (20.6–21.1 Å). However, the absolute abundance is a ratio of the line fluxes to the continuum fluxes, and this is less well determined in the RGS data (compared to the EPIC data) for two reasons. The first is that the S/N is poorer in the RGS data due to fewer photons. The second point is that there is a power-law continuum contribution from LMXBs, which can be constrained at high energies from the higher S/N EPIC data. This power-law component cannot be properly constrained with the RGS data because of lower S/N and because the instrument has no response above 2.5 keV, where the power law is most distinct. Therefore, we found it useful to consider the EPIC and RGS spectral fits separately as...
well as jointly. All spectral fitting was performed within XSPEC (ver. 11.3.2).

As discussed below, there is good consistency when fitting the EPIC instruments separately, so we perform a simultaneous fit of a single model to the MOS1, MOS2, and the PN data with their normalizations free to fit. For the thermal emission from thermal gas, we used the Astrophysical Plasma Emission code (APEC; Smith et al. 2001), which models optically thin emission from an isothermal plasma of some abundances (vapec). The optically thin assumption should be appropriate in most cases, although Xu et al. (2002) demonstrate that optical depth effects occur in the strongest Fe xvi lines for the galaxy NGC 4636. Variations in temperature are accommodated by introducing two vapec models with different temperatures but the same set of abundances. For the fits, we fix the X-ray absorbing column at the Galactic H i value (Dickey & Lockman 1990), and we fix for the temperature, the abundances of Fe, O, Si, and Mg, and the normalization. When the data had sufficiently high S/N, we also fit for Ne, Ni, and S; otherwise, they were tied to Fe with the ratios given by the solar values. All other elemental abundances were tied to Fe, except for He, which was set at the solar value. The abundances are referenced to the solar values, which are defined here by the values of Grevesse & Sauval (1998). In many earlier works, the solar abundances of Anders & Grevesse (1989) were used, which have higher concentrations of O and Fe by 26% and 48%. Using the newer solar abundances, the relative abundances are raised.

When fitting for the power-law component, most point sources are unresolved, unlike the situation for the Chandra observations. Based on Chandra observations of point sources in early-type galaxies, the summed X-ray point source spectra can be represented by a power-law model with an average photon index Γ of 1.56 (Irwin et al. 2002). This component dominates mainly the hard X-ray band (> 2 keV), but also contributes some emission to the soft X-ray band (< 2 keV). In the models, the contribution from this model is implemented by including a power-law component, with a fixed photon index of Γ = 1.56 and a free normalization; this component adds one free parameter. From our tests, we find nearly the same value for the power-law normalization when we only consider data above 3 keV (and no thermal component) or when we fit the entire energy range of the spectrum, using a power-law and a vapec component.

The procedure for fitting the Chandra spectra are similar except that most of the point sources are excluded from the spectral fits. Still, there remains the contribution from unresolved point sources, so we included a power-law component, with fixed Γ to 1.56, along with the thermal component (vapec) and a Galactic absorption term (wabs).

For the fits to the RGS data, there are a few other considerations, the most important being the broadening of the lines along the dispersion direction due to the extended nature of the emission. A correction for this can be made if the spatial distribution of the emission is known, and for this purpose, we used the EPIC MOS1 image (equivalently, the other images could have been used). Using this image, the broadening correction is implemented by the task rgsxspec within XSPEC. In the first set of fits to the RGS data, we fix the ratio of the luminosities from the thermal plasma component (vapec) to the power-law component (powerlaw) at the value given by the fits to the EPIC data. For the thermal plasma model, we use the RGS data to fit for the temperature, the normalization, and the abundances for O and Fe, which are the element abundances that are best constrained by the RGS data. Abundances for Mg and Si are determined whenever possible, otherwise they are tied to Fe. In addition, lines are detected from Ne and N for some sources, but the S/N is not sufficient for useful determinations, so the abundances of N, Ne, S, and Ni are tied to Fe.

### 3.2. Comparison of ACIS-S and EPIC Spectral Fits for Bright Sources

In the comparison between the four imaging devices (see Figure 2), the results for the inner 1’ diameter of NGC 4649 were fairly consistent, especially between the different EPIC instruments. Using the model described above, best fits were obtained separately for the ACIS-S3 data, the PN data, the combined MOS1+MOS2 data, and the combined RGS1+RGS2 data. This galaxy is an important case because the temperature gradient is small and there is no evidence for a multi-temperature gaseous medium. We find that the temperatures were nearly identical for all detectors, but there were differences in the abundances of O and S between the XMM-Newton instruments and the Chandra ACIS-S detector. The S abundance obtained from the ACIS data is 75% lower than that derived from the EPIC data, and the 90% error bars do not overlap. However, the two abundance measurements differ at the 1.8σ level, so this is not a major discrepancy. A similar, but slightly greater discrepancy exist for the O abundance, where the EPIC value is about 2σ above the ACIS value. Aside from these differences, the abundances follow rather closely for the different instruments in that, relative to solar values, O is the least abundant element, followed by Ne. Relative to oxygen, iron has a significantly higher abundance, with Mg and Si a bit higher yet, and then Ni is much higher than the other elements.
3.3. Comparison of the RGS and EPIC Results

We considered one-temperature (1T) models in both the central 1′ (for both EPIC and RGS data) and more extended 4′ diameter regions (for EPIC)/3.5 regions (for RGS), as well as 2T models, shown in Tables 2–3. For the single-temperature fits (Table 2), we give both the EPIC (MOS + PN) and the RGS results, and aside from the temperature and abundances, we list the degrees of freedom in the fit (dof), the reduced $\chi^2$, and the luminosities from the power-law component (Power) and from the thermal component in the 0.33–2.48 keV band (in units of $10^{39}$ erg s$^{-1}$). The final column is the ratio of the thermal to the power-law luminosities. For the fit to the RGS data, we forced this ratio to be the same value as for the EPIC fit. In Table 3, we fit all the EPIC data and some good quality RGS data, but with a 2T thermal model, so there are two additional columns, for the second temperature, and for the luminosity from the second thermal component. The ratio of the luminosities is the sum of the two thermal components divided by the power-law component. There are two rows per instrument in each table, for the 1′ (for both EPIC and RGS) and the 4′ (for EPIC)/3.5 (for RGS) diameter regions.

When fitting a model to the data, only photon statistics define the standard deviation per point. In addition, there are systematic uncertainties in the calibration of the instrument as well as instrumental artifacts, not all of which are known or excluded from spectral fits. There is no way of including accurately such effects in the spectral fitting, but we can identify the way in which the $\chi^2$ will be affected. The resulting $\chi^2$ will be larger than if there were no systematic effects. Furthermore, these systematic effects do not change as the source becomes brighter, but brighter sources have more photons and therefore lower uncertainties due to photon statistics. Therefore systematic errors in valid spectral fits will lead to significantly higher reduced $\chi^2$ and fits that formally would be rejected. This can lead to one rejecting spectral fits to all of the best data sets. Consequently, we have inspected every spectral fit and there are several that appear to be good fits but would formally be rejected. This can be seen in the large image size, since usually the isolated O lines are large. Those residuals are not due to improper convolution with the instrumental profile. The fit using the 4′ diameter region is better than the central 1′ region results.

In these 1T models, there is quite good agreement in the value of the temperature between the EPIC and RGS fits, and there is also good agreement in the Fe and O abundances between the EPIC and RGS fits with within 90% error range except for NGC 720, NGC 3923, NGC 4552, and IC 1459, which have very poor RGS spectra to constrain abundance parameters (Table 2). Excluding the galaxies NGC 1399, NGC 4472, NGC 4636, and IC 1459 (see above), the range of the Fe and O abundances from the EPIC are 0.42–1.57 and 0.26–0.55, respectively (relative to the solar values). The median values for Fe and O are 0.58 and 0.35 of the solar values, while the median values from the RGS fits (also excluding NGC 720, NGC 3923, NGC 4552) are slightly higher, at 0.65 and 0.42. The ratio of the Fe to O abundance is around 2:1 in this data set.

Of particular interest are three galaxies, NGC 4406, NGC 4649, and NGC 5044, where the RGS fits appear to be good and where the abundance uncertainties are not large (despite having formally unacceptable values of $\chi^2$, discussed above). The galaxy NGC 4649 was discussed above and here we add that the RGS metallicities are in good agreement with those derived from the EPIC data. This is one of the few galaxies where the Fe metallicity is supersolar, although only modestly. The galaxy NGC 5044 has consistent abundances from the EPIC and RGS data, which show that the Fe abundance is about 0.74 of solar and the O abundance is about 0.43 of solar. The third galaxy, NGC 4406, has consistent O abundances (about 0.36 solar) between the two XMM-Newton instruments, but the Fe abundance is higher in the EPIC observations. This is a complicated galaxy that is interacting strongly with its environment (Stickel et al. 2003), so perhaps it is not surprising that there would be disagreement for two instruments that sample slightly different spatial regions.

When a second thermal component was introduced, the spectral fits generally improved and the metallicities usually increased. (We fit only six sources with good quality RGS spectra for a second thermal component added. The second temperature will be frozen to its EPIC value if it can not be constrained. We list all results but NGC 4649 in Table 3, which shows no convergence for the 2T model.) For the same eight galaxies used here (all except NGC 4636 and IC 1459) and for 1T fits, the Fe abundance rose by a median of 30% while the O abundance rose by 4%. For the EPIC fits to the data within the central 1′ (Table 3), four of the galaxies had formally acceptable fits (NGC 720, NGC 3923, NGC 4552, and NGC 5044), NGC 4406 is nearly acceptable, and the fits to the others appear to be good, with the exception of NGC 4636. Of the good and formally acceptable fits (all except NGC 4636 and IC 1459), the range in the Fe abundance is 0.44–1.70 solar and for O it is 0.27–0.76 solar, with medians of 0.86 and 0.44 solar. Even the two galaxies excluded from the analysis had similar values. The abundances for Mg and Si are indistinguishable from that of Fe, whereas in the 1T fits, they were systematically higher than the Fe abundance. In these fits, most of the power is from the gas phase, typically more than 90% in the 0.33–2.48 keV band. The RGS fits have Fe and O abundances of 0.33–1.50 solar and 0.31–1.22 solar, with medians of 0.95 and 0.66 solar, which is consistent in Fe but higher in O.

In Figures 4–9, we show zoomed-in RGS spectra with best fit model for the six most bright galaxies with marked individual emission lines. Note that, overall the fits are fairly good, but there are some residuals at Fe xvii and Fe xviii at 15 Å, 16 Å, or 17 Å. Those residuals are not due to improper convolution with large image size, since usually the isolated O viii line at 19 Å is fitted fairly well. They might be due to calibration issues or non-accurate APEC model at those lines.

One important object for comparison is NGC 4636, where Xu et al. (2002) fit a model to the RGS data within 2′ that included a
| Galaxy      | Camera/Size | T(k)eV | Fe    | O     | Mg    | Si     | Ne     | S     | Ni     | dof | χ²e | Power | Vape | Vape/Power |
|-------------|-------------|--------|-------|-------|-------|--------|--------|-------|--------|-----|-----|-------|------|------------|
| NGC 720     | EPIC 1’     | 0.560  | 0.51  | 0.30  | 0.40  | 0.35   | 0.54   | ...   | ...    | 198 | 1.26 | 0.54  | 1.32 | 3.22       |
|             | EPIC 4’     | 0.396  | 0.52  | 0.49  | 0.45  | 0.35   | 0.54   | ...    | ...    | 500 | 0.79 | 1.61  | 4.19 | 2.61       |
|             | RGS 1’      | 0.369  | 1.54  | 0.10  | 0.07  | 0.06   | 0.06   | ...    | ...    | 62  | 1.31 | 0.92  | 2.87 | 3.05       |
|             | RGS 3.5’    | 0.327  | 4.10  | 0.35  | 0.22  | ...    | ...    | ...    | ...    | 143 | 1.30 | 2.06  | 5.37 | 2.61       |
| NGC 1399    | ACIS-S 1’   | 0.809  | 1.33  | 0.77  | 1.49  | 0.58   | 0.76   | ...    | ...    | 143 | 3.52 | 0.71  | 7.98 | 11.23      |
|             | EPIC 1’     | 0.813  | 1.25  | 0.54  | 1.65  | 1.40   | 1.97   | 1.13   | ...    | 775 | 2.39 | 0.63  | 7.68 | 12.18      |
|             | EPIC 4’     | 0.950  | 0.65  | 0.89  | 1.06  | 0.89   | 0.89   | ...    | ...    | 1367| 2.93 | 3.46  | 19.94| 5.77       |
|             | RGS 1’      | 0.813  | 1.32  | 1.20  | 0.36  | ...    | ...    | ...    | ...    | 1011 | 1.19 | 0.86  | 14.12| 16.46      |
|             | RGS 3.5’    | 0.994  | 0.68  | 0.55  | 0.44  | 0.98   | 0.44   | ...    | ...    | 1451 | 1.32 | 4.55  | 26.88| 5.91       |
| NGC 3923    | EPIC 1’     | 0.479  | 0.49  | 0.26  | 0.63  | 0.27   | 0.27   | ...    | ...    | 315 | 1.34 | 0.61  | 3.48 | 5.68       |
|             | EPIC 4’     | 0.387  | 0.63  | 0.60  | 0.96  | 0.58   | 0.58   | ...    | ...    | 588 | 1.16 | 1.86  | 5.20 | 2.80       |
|             | RGS 1’      | 0.446  | 5.00  | 2.06  | 0.69  | 0.83  | ...    | ...    | ...    | 9   | 1.42 | 0.64  | 3.76 | 5.86       |
|             | RGS 3.5’    | 0.432  | 5.00  | 3.01  | 1.68  | ...    | ...    | ...    | ...    | 184 | 1.47 | 1.44  | 4.66 | 3.24       |
| NGC 4406    | EPIC 1’     | 0.647  | 0.64  | 0.40  | 0.65  | 0.69  | 0.83  | ...    | ...    | 462 | 1.27 | 0.01  | 0.08 | 8.26       |
|             | EPIC 4’     | 0.716  | 0.65  | 0.49  | 0.65  | 0.65  | 0.82  | 0.56   | ...    | 919 | 1.40 | 0.04  | 0.44 | 12.07      |
|             | RGS 1’      | 0.626  | 0.34  | 0.32  | 0.15  | ...    | ...    | ...    | ...    | 304 | 1.29 | 0.03  | 0.29 | 8.32       |
|             | RGS 3.5’    | 0.623  | 0.29  | 0.25  | 0.16  | ...    | ...    | ...    | ...    | 774 | 1.23 | 0.06  | 0.80 | 13.29      |
| NGC 4472    | ACIS-S 1’   | 0.724  | 1.58  | 0.71  | 1.99  | 1.87  | 2.15  | 1.00  | 3.35   | 123 | 1.59 | 0.29  | 4.15 | 14.11      |
|             | EPIC 1’     | 0.748  | 1.11  | 0.54  | 1.17  | 1.19  | 0.83  | 0.99   | 3.93   | 720 | 1.77 | 0.40  | 4.36 | 11.01      |
|             | EPIC 4’     | 0.813  | 1.05  | 0.51  | 1.32  | 2.81  | 1.18  | 5.36   | ...    | 1162| 5.23 | 1.38  | 10.50| 7.59       |
|             | RGS 1’      | 0.752  | 0.71  | 0.59  | 0.53  | ...    | ...    | ...    | ...    | 536 | 1.47 | 0.56  | 8.37 | 15.01      |
|             | RGS 3.5’    | 0.812  | 0.10  | 0.91  | 0.81  | ...    | ...    | ...    | ...    | 1015| 1.23 | 1.50  | 13.56| 9.06       |
| NGC 4552    | EPIC 1’     | 0.608  | 0.42  | 0.29  | 0.46  | 0.40  | ...    | ...    | ...    | 414 | 1.06 | 0.07  | 0.18 | 2.64       |
|             | EPIC 4’     | 0.386  | 0.66  | 0.22  | 1.45  | 1.16  | ...    | ...    | ...    | 628 | 1.29 | 0.13  | 0.22 | 1.69       |
|             | RGS 1’      | 0.557  | 4.32  | 4.61  | 4.44  | ...    | ...    | ...    | ...    | 89  | 1.64 | 0.07  | 0.18 | 2.59       |
|             | RGS 3.5’    | 0.376  | 5.00  | 2.59  | ...    | ...    | ...    | ...    | ...    | 190 | 1.46 | 0.15  | 0.25 | 1.67       |
| NGC 4636    | EPIC 1’     | 0.535  | 0.62  | 0.41  | 0.61  | 0.24  | ...    | ...    | ...    | 506 | 2.06 | 0.14  | 4.83 | 34.70      |
|             | EPIC 4’     | 0.631  | 0.92  | 0.58  | 0.69  | 0.74  | ...    | ...    | ...    | 838 | 2.18 | 0.53  | 14.89| 27.92      |
|             | RGS 1’      | 0.622  | 0.83  | 0.38  | 0.45  | ...    | ...    | ...    | ...    | 438 | 2.02 | 0.20  | 6.93 | 34.92      |
|             | RGS 3.5’    | 0.639  | 0.91  | 0.43  | 0.48  | ...    | ...    | ...    | ...    | 893 | 1.70 | 0.52  | 14.57| 27.91      |
|             | RGS corrected 1°| 0.589  | 1.00  | 0.53  | 0.65  | ...    | ...    | ...    | ...    | 389 | 1.34 | 0.20  | 7.19 | 36.22      |

Table 2: One-Phase Model Fitting for Chandra ACIS-S, XMM-Newton EPIC, and RGS Spectra
## Table 2
(Continued)

| Galaxy     | Camera/Size$^a$ | $T$(keV)$^b$ | Fe $^c$ | O $^c$ | Mg$^c$ | S$^c$ | Ne$^c$ | S$^c$ | Ni$^c$ | dof$^d$ | $\chi^2$ | Power$^f$ | Vpec$^g$ | Vpec/Power$^h$ |
|------------|----------------|-------------|---------|--------|--------|--------|--------|--------|--------|--------|----------|----------|----------|
| RGS corrected 3/5$^i$ | 0.625$^{+0.007}_{-0.006}$ | 1.07$^{+0.11}_{-0.02}$ | 0.57$^{+0.03}_{-0.07}$ | 0.65$^{+0.12}_{-0.17}$ | - | - | - | - | 810 | 1.307 | 0.52 | 14.85 | 28.48 |
| NGC 4649   |                |             |         |        |        |        |        |        |        |        |          |          |          |
| ACIS-S 1$^j$ | 0.763$^{+0.014}_{-0.017}$ | 1.39$^{+0.21}_{-0.19}$ | 0.13$^{+0.17}_{-0.13}$ | 1.46$^{+0.29}_{-0.28}$ | 1.47$^{+0.29}_{-0.27}$ | 0.00$^{+0.55}_{-0.00}$ | 0.31$^{+0.19}_{-0.31}$ | - | - | 812 | 1.421 | 0.48 | 5.68 | 11.91 |
| EPIC 1$^j$  | 0.776$^{+0.003}_{-0.003}$ | 1.57$^{+0.12}_{-0.11}$ | 0.55$^{+0.11}_{-0.10}$ | 1.62$^{+0.18}_{-0.16}$ | 1.63$^{+0.17}_{-0.15}$ | 0.77$^{+0.12}_{-0.24}$ | 1.25$^{+0.20}_{-0.20}$ | 4.30$^{+0.45}_{-0.43}$ | 630 | 1.408 | 0.58 | 5.66 | 9.73 |
| EPIC 4$^j$  | 0.784$^{+0.002}_{-0.002}$ | 1.39$^{+0.08}_{-0.08}$ | 0.58$^{+0.08}_{-0.07}$ | 1.43$^{+0.12}_{-0.11}$ | 1.64$^{+0.12}_{-0.11}$ | 0.70$^{+0.16}_{-0.17}$ | 1.09$^{+0.15}_{-0.15}$ | 3.63$^{+0.16}_{-0.30}$ | 931 | 1.506 | 1.75 | 9.63 | 5.50 |
| RGS 1$^j$  | 0.759$^{+0.010}_{-0.010}$ | 1.37$^{+0.55}_{-0.20}$ | 0.67$^{+0.33}_{-0.07}$ | - | - | - | - | - | - | 325 | 1.335 | 0.71 | 6.93 | 9.73 |
| RGS 3/5    | 0.776$^{+0.014}_{-0.014}$ | 1.53$^{+0.18}_{-0.14}$ | 0.96$^{+0.21}_{-0.36}$ | 1.04$^{+0.32}_{-0.42}$ | - | - | - | - | 552 | 1.258 | 1.66 | 9.41 | 5.67 |
| NGC 5044   |                |             |         |        |        |        |        |        |        |        |          |          |          |
| EPIC 1$^j$  | 0.751$^{+0.002}_{-0.005}$ | 0.82$^{+0.06}_{-0.06}$ | 0.44$^{+0.09}_{-0.08}$ | 0.80$^{+0.11}_{-0.10}$ | 0.80$^{+0.10}_{-0.09}$ | 0.44$^{+0.18}_{-0.18}$ | 0.71$^{+0.15}_{-0.15}$ | 2.24$^{+0.37}_{-0.36}$ | 477 | 1.118 | 1.83 | 49.05 | 26.87 |
| EPIC 4$^j$  | 0.792$^{+0.002}_{-0.002}$ | 0.72$^{+0.02}_{-0.02}$ | 0.37$^{+0.04}_{-0.04}$ | 0.94$^{+0.05}_{-0.05}$ | 0.89$^{+0.04}_{-0.04}$ | 0.56$^{+0.08}_{-0.08}$ | - | - | 832 | 1.963 | 8.80 | 249.84 | 28.39 |
| RGS 1$^j$  | 0.771$^{+0.010}_{-0.011}$ | 0.65$^{+0.09}_{-0.07}$ | 0.42$^{+0.10}_{-0.08}$ | 0.89$^{+0.33}_{-0.30}$ | - | - | - | - | 295 | 1.262 | 4.60 | 123.97 | 26.95 |
| RGS 3/5    | 0.792$^{+0.007}_{-0.007}$ | 0.59$^{+0.03}_{-0.04}$ | 0.46$^{+0.07}_{-0.06}$ | 0.49$^{+0.14}_{-0.19}$ | 0.85$^{+0.19}_{-0.19}$ | - | - | - | 668 | 1.210 | 10.25 | 293.81 | 28.65 |
| IC 1459    |                |             |         |        |        |        |        |        |        |        |          |          |          |
| EPIC 1 no core$^i$ | 0.612$^{+0.036}_{-0.028}$ | 1.62$^{+0.25}_{-0.25}$ | - | - | - | - | - | - | 171 | 0.828 | 0.90 | 0.31 | 0.35 |
| EPIC 4 no core$^i$ | 0.583$^{+0.020}_{-0.020}$ | 0.39$^{+0.14}_{-0.09}$ | - | - | - | - | - | - | 462 | 1.049 | 1.96 | 1.36 | 0.69 |
| RGS 1$^i$  | 0.532$^{+0.063}_{-0.089}$ | 1.18$^{+3.40}_{-0.84}$ | 2.32$^{+2.68}_{-1.79}$ | - | - | - | - | - | 66 | 1.105 | 2.02 | 1.95 | 0.97 |
| RGS 3/5    | 0.307$^{+0.114}_{-0.012}$ | 4.96$^{+0.04}_{-0.06}$ | 1.84$^{+3.16}_{-1.64}$ | - | - | - | - | - | 138 | 0.994 | 2.86 | 2.82 | 0.99 |

Notes.

$^a$ Instruments used and their apertures. EPIC: joint fit for MOS1,MOS2,& pn data; RGS: joint fit for RGS1 & RGS2 data. EPIC source size: diameters of 1’ & 4’; RGS source size: including 90% and 99% of PSF, around 1’ and in the range of 0.33 keV to 2.48 keV.

$^b$ One-phase temperature in unit of keV.

$^c$ Elemental abundances of Fe, O, Mg, Si, Ne, S, and Ni relative to their solar values (Grevesse & Sauval 1998). Note that for RGS data, the Mg and Si abundances were allowed to fit when thawing them can improve the fits, otherwise they were tied to Fe.

$^d$ Degrees of freedom.

$^e$ Reduced $\chi^2$ of the fit.

$^f$ Luminosity of the power-law component in unit of $10^{46}$ erg s$^{-1}$ and in the range of 0.33 keV to 2.48 keV.

$^g$ Luminosity of the thermal component (vpec) in unit of $10^{46}$ erg s$^{-1}$ and in the range of 0.33 keV to 2.48 keV.

$^h$ Luminosity ratio of the thermal component to the power-law one.

$^i$ We excluded the Fe xvi line at 15 Å (i.e. excluding 14.75–15.40 Å) to correct the optical depth effect.

$^j$ We excluded the central 13” radius region to eliminate the central non-hot-gas emission.
### Table 3

Two-Phase Model Fitting for XMM-Newton EPIC and RGS Spectra

| Galaxy    | Camera/Size | $T_1$(keV) | $T_2$(keV) | Fe% | O% | Mg% | Si% | Ne% | S% | Ni% | dof | $\chi^2$/dof | Power | Vape1 | Vape2 | Vape3/Power |
|-----------|-------------|------------|------------|-----|----|-----|-----|-----|----|-----|-----|------------|--------|--------|--------|------------|
| NGC 720 EPIC 1' | $0.219^{+0.161}_{-0.109}$ | $0.567^{+0.019}_{-0.017}$ | $0.64^{+0.38}_{-0.17}$ | $0.36^{+0.24}_{-0.12}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ | $0.56^{+0.46}_{-0.23}$ | $0.49^{+0.47}_{-0.33}$ |

**Notes.** Table is the same as previous one, only that we use two-phase model here: \( wabd(\text{vape}1+\text{vape}2+\text{power}) \) for EPIC, and \( wabd(\text{vape}1+\text{vape}2+\text{power}^*\text{rgxsrec}) \) for RGS. For RGS spectra, the second temperature will be frozen to the EPIC fits if it cannot be constrained.
temperature gradient and optical depth effects for Fe \textsuperscript{xvii}, which necessitated the use of a $\beta$ model. Their abundances for Fe and O are $1.29 \pm 0.06$ and $0.63 \pm 0.06$, relative to the solar abundances of Grevesse & Sauval (1998). Their ratio of Fe to O is 2 in these units. From our fitting of the RGS data with a single temperature and no opacity correction, we obtain Fe and O abundances that are about 30% lower, but with the same abundance ratio. Our somewhat lower abundance for Fe might be due to the absence of optical depth corrections. Xu et al. (2002) found that the optical depth of Fe \textsuperscript{xvii} line at 15.0 Å is greater than unity for the densities and temperatures in the core of NGC4636 ISM, while the Fe \textsuperscript{xvii} blend at 17.1 Å is negligible. Many scatterings of the 15.0 Å photons flatten its profile comparing to the others, which would cause large uncertainty on Fe abundance. To correct the optical depth effect at 15.0 Å, we fit again the 3:5 aperture RGS spectra of those very bright galaxies (NGC1399, NGC 4406, NGC 4472, NGC4636, NGC 4649, and NGC 5044), excluding the problematic 15.0 Å line (i.e. excluding 14.75–15.4 Å) as well as freeing the redshift parameter to account any residual line shift. It turns out that only NGC 4636 shows large changes in both parameter values (18% increase in Fe, and 33% increase in O) and in reduced $\chi^2_r$ (decrease by 23%). Other galaxies only show at most 4% variances in Fe, O, and reduced $\chi^2_r$. So the optical depth effect at the 15.0 Å line is negligible for all

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**Figure 3.** Example of NGC 4649 for XMM-Newton EPIC spectral fitting with marked emission lines. The upper panel is for 1′ aperture region while the lower panel is for 4′ aperture region. (Green): pn camera; (red): MOS1 camera; (black): MOS2 camera. Note: the emission lines from Fe L complex and O \textsuperscript{viii} are completely unresolved while emission lines from Mg, Si, and S can be identified fairly well in CCD spectra.

(A color version of this figure is available in the online journal.)
Figure 4. The 10–20 Å region of the XMM-Newton RGS1 and RGS2 spectra for NGC 1399 with the best-fit models. The upper panel is for 1’ aperture while the bottom one is for 3.5’ aperture, with marked resolved emission lines mostly from the Fe L complex and O viii. Note that the wide gap between 10.5 Å and 13.8 Å is due to failure of RGS1 CCD7, while the narrower gaps at 11.4 Å, 13.1 Å, 14.6 Å, 15.3 Å, 16.5 Å, 17.1 Å, and 18.9 Å are due to gaps between RGS CCDs.

3.4. Comparison with Other Studies of the Metal Abundances in Early-type Galaxies

We show below the detailed comparison with other studies of the metal abundances in early-type galaxies for individual galaxy. Note that our fitted temperatures agree well with the published results, and abundances from literatures are scaled to Grevesse & Sauval (1998).

NGC 720. Our 4’ diameter aperture EPIC results for NGC 720 (Fe(1.10±0.22), O(0.45±0.14), Mg(0.76±0.32), and Si(0.97±0.44)) are consistent with Chandra results reported by Humphrey & Buote (2006) (Fe(0.71±0.21), O(0.16±0.16), and Mg(0.90±0.37)), and are also consistent with the 180 ks Suzaku observation by Tawara et al. (2008) within 90% error range, who obtained abundances of Fe(0.73±0.11), O(0.47±0.10), Mg(0.56±0.10), and Si(0.54±0.18).

NGC 1399. This is an important galaxy with supersolar abundances reported by previous observations. Buote (1999) found an Fe abundance of 2.40±1.06 within a 10’ diameter aperture by ASCA. Later, he reported an XMM-Newton result for this galaxy that the Fe abundance in the central region of the galaxies except for NGC 4636. After correcting this effect for NGC 4636, we still have slightly lower abundances of Fe (1.07±0.11) and O (0.57±0.07) from 1T fit, but have consistent abundances of Fe (1.42±0.16) and O (0.58±0.05) from the 2T fit, compared to Fe (1.29 ± 0.06) and O (0.63 ± 0.06) from Xu et al. (2002). These abundances are slightly higher than what XMM-Newton EPIC data and ASCA data reveal. Buote (1999) obtained abundances of Fe and O with ASCA of 1.08±0.24 and 0.49±0.29, respectively, which are consistent with our 2T fit for XMM-Newton EPIC data at 4’ diameter region, 1.06 ± 0.03 for Fe and 0.51±0.03 for O. The reason that these abundances from CCD spectral analysis are systematically lower (25% lower for Fe and 12% lower for O) than the XMM-Newton RGS results is not due to the uncorrected optical depth effect for Fe xvi line at 15 Å, since we refitted the XMM-Newton EPIC 4’ diameter region excluding this line (i.e. excluding 0.79–0.88 keV), and found only 4% increase in Fe and 2% increase in O. The reason might be that the RGS and EPIC cameras cover different regions of this galaxy, which shows complex structure and events that have been described by others (Jones et al. 2002; O’Sullivan et al. 2005).
(about 10′ in diameter) is 1.54–2.04 (Buote 2002). Our net 120 ks XMM-Newton observation also shows a consistent but more constrained Fe abundance, 1.32±0.05, within 4′ diameter region. This Fe abundance is higher than what Humphrey & Buote (2006) obtained with Chandra data, 1.06 ± 0.09. This is because Humphrey & Buote (2006) obtained an emission-weighted average Fe abundance up to 20′ in diameter, which tends to be lower because the Fe abundance decreases to subsolar at large radius (20′ in diameter) of NGC 1399 (Humphrey & Buote 2006).

**NGC 3923.** Humphrey & Buote (2006) reported a poorly constrained abundance of Fe for 1.03(>0.24) due to very few counts of photons. We obtained a lower but more constrained Fe abundance (0.72±0.25). Since we have many more counts of photons (total counts of 3.3 × 10^4) for this galaxy and obtained a good fit ($\chi^2_r = 1.024$), we are confident with our result.

**NGC 4406.** ASCA observation shows an Fe abundance of 0.77+0.33 (Buote et al. 1998), while Chandra observation shows Fe and O abundances of 1.10±0.47 and 0.46+0.13, respectively (Athey 2007). Our results are consistent with these two within error range, which are Fe of 0.87+0.02−0.04 and O of 0.51+0.05−0.05.

**NGC 4472.** Buote (1999) reported an ASCA result for this galaxy to be 2.96+3.55−1.14 for Fe abundance, while Humphrey & Buote (2006) reported an emission-weighted average abundances for Fe and O of 1.25+1.52−0.36 and 0.48+0.18−0.12, respectively. Our XMM-Newton result gives consistent but more constrained Fe and O abundances of 1.75+0.07−0.05 and 0.77+0.06−0.05, respectively.

**NGC 4552.** Humphrey & Buote (2006) measured abundances of Fe and O from Chandra data of 0.63+0.04−0.03 and 0.08 ± 0.06, while Athey (2007) reported abundances also measured from Chandra data up to 2.5 diameter region to be 0.73+0.07−0.06 for Fe and 0.29+0.04−0.18 for O. We obtained better constrained abundances of Fe and O of 0.49±0.08 and 0.25+0.07−0.18, respectively, which are consistent with Athey (2007)’s Fe and O abundances within 90% error range, but is 68% higher in O abundance than Humphrey & Buote (2006)’s result. We attribute this discrepancy to be not having enough counts in Chandra’s spectra to constrain the O abundance.

*Extra comment for NGC4552.* This system is known to be suffering fairly strong ram-pressure stripping and the extent of its gas is quite well defined from the Chandra images (Machacek et al. 2006). The 1′ diameter region is quite a good match to the gas, but the 4′ aperture might contain a large fraction of

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**Figure 5.** Same as Figure 4, but for NGC 4406.

(A color version of this figure is available in the online journal.)
cluster and mixed cluster/galaxy gas. In our two-phase model fits in Table 3, however, the abundances are quite consistent between 1' and 4' regions. Therefore, the cluster emission will not significantly affect the measurement of the galaxy hot gas abundances in NGC4552 even for 4' aperture. The reason may be that even for the 4' region, the local hot gas in NGC4552 is still the dominant thermal component comparing to other thermal component (like the cluster emission), which contributes 63% of the total thermal luminosity.

**NGC 4636.** See detailed discussion in Section 3.3.

**NGC 4649.** See detailed discussion in Section 3.2.

**NGC 5044.** Tamura et al. (2003) obtained abundances of Fe and O within the central 10–20 kpc region (2' width) of NGC 5044 from XMM-Newton RGS data. Their abundances for Fe and O are 0.81 ± 0.07 and 0.31 ± 0.13, respectively, which is consistent in O with but higher in Fe than our RGS 1T model-fitting results for 1' region, Fe(0.65±0.05), O(0.42±0.10), but are consistent with our RGS 2T model-fitting results of Fe(0.77±0.12) and O(0.48±0.12). Athey (2007) used Chandra for annualar analysis for this galaxy, and found abundances of Fe, O, Mg, and Si within 4' diameter in the ranges of 0.92–1.53, 0.18–0.96, 1.01–1.77, and 0.95–1.93, respectively. These abundances gradients are also conformed by our study with two aperture analysis of EPIC data. We find abundances gradients of Fe, O, Mg, and Si between smaller and larger apertures in the ranges of 0.86–1.15, 0.41–0.49, 0.84–1.14, and 0.86–1.01, respectively. Our average Fe abundance (1.15 ± 0.05) within 4' diameter is also consistent with the result by Buote et al. (2003), who reported the XMM-Newton observation of emission-weighted average abundance of Fe within 10' diameter to be 1.09 ± 0.04.

**IC 1459.** This is the only galaxy in our sample that shows different spectral characteristics due to a strong central point source (Fabbiano et al. 2003). Its power-law component is extremely large compared to the thermal component, making it hard to constrain the abundances in the thermal emission. We fit its spectra excluding the central 13'' region in radius, which is mainly a power-law component with a photon index $\Gamma$ of 1.88 and an excessive absorption $n_H$ of $2.9 \times 10^{21}$ cm$^{-2}$ (Fabbiano et al. 2003). We then fit the remainder (mainly thermal emission) with a 1T thermal component plus a power-law (photon index $\Gamma$ fixed to 1.56) with only Fe allowed to vary (other elements are tied to Fe). This gives a better constrain on the Fe abundance of 0.39±0.09, which is consistent with the Chandra result of 0.74$^{+0.61}_{-0.44}$ (Athey 2007) and the ASCA result of 0.21$^{+0.09}_{-0.13}$ (Buote et al. 1998), within their error ranges. Longer observation is needed to constrain other element abundances.
To test if the central point source contributes any photons outside the excluded 13″ region, we refitted the IC1459 EPIC 1′ region excluding the central 13″ core with photon index free to fit. The fitted photon index is 1.60$^{+0.06}_{-0.05}$, only increased by 2.6% from 1.56, and the temperature and Fe abundance change not much, by 0.5% and 3.7%, respectively. For the 4′ region, the results are essentially not changed. So we conclude that the central point source in IC1459 does not contribute many photons outside the excluded 13″ region.

3.5. Discussion on the Assumptions of the Spectral Fitting Method

In our spectral fitting method, we made two simplifying assumptions in the choice of spectral models: freezing the slope of the power-law component and linking the abundances of the 2T models. We did the following tests to justify these assumptions.

Effect on the slope of the power-law component. To check the effect on photon index, we test two cases (with 2T model plus a power law), NGC5044 and NGC3923, according to their different fractions of the total luminosity for the power-law component. During the fitting, the photon index is free to fit. For NGC5044, the power-law component only contributes 3.4% to the total luminosity in the energy band of 0.33 keV to 2.48 keV in its central 1′ diameter region. The fitted photon index is 1.80$^{+0.27}_{-0.28}$, which is consistent with 1.56 within its 90% error range. The two temperatures vary by only 0.3%, while the abundances vary at most 5.8%, which are within 90% confidence error range of previous result when fixed photon index to 1.56.

For NGC3923, the power-law component contributes 24.6% to the total luminosity in the energy band of 0.33 keV to 2.48 keV in its central 4′ diameter region. The fitted photon index is 1.73$^{+0.09}_{-0.20}$, which is consistent with 1.56 within its 90% error range. There are no significant changes in reduced chi-square value as well as the two temperature values, only by 0.4%. Abundance values, however, show large increase by at least 50%. Fe changes most, increased by 70.8%, from 0.72$^{+0.25}_{-0.19}$ to 1.23$^{+1.07}_{-0.44}$, but is still consistent within their error range.

So we conclude that the value of photon index does not significantly affect the abundance values for hot gas dominant
galaxies, and freezing photon index to 1.56 can better constrain the abundance values. But it is the main error source for galaxies where power-law component is comparable to the hot gas component. In our sample, most sources only have a few percent of the total luminosity for the power-law component, and thus are not greatly affected by the value of photon index. Only four galaxies (NGC720, NGC3923, NGC4552, and IC1459) have power-law component with luminosity greater than 15% of the total luminosity, thus their abundances are more greatly affected by the value of photon index. Longer observations for these sources are needed to better constrain the abundances.

**Effect on linking the abundances of the 2T models.** There are five galaxies (NGC1399, NGC4406, NGC4472, NGC4649, and NGC5044) in our sample which show the possible signs of overlying cluster emission. This can be seen from the 2T model fitting temperatures (see Table 3), where the galaxy hot gas temperatures are around 0.63–0.81 kev, while the cluster emission temperatures are around 0.94–1.72 kev.

We chose NGC1399 4′ region to test the effect of linking two abundance sets, because it has the most photon counts and a comparable cluster emission component to the local galaxy hot gas (about 46% of the total thermal luminosity). During the fit, we unlinked the two abundance sets. Each abundance set then has free parameters of Fe, O, Mg, Si, S, Ne, and Ni. Other elements are tied to Fe. The fitting results show no change for the reduced $\chi^2$ value, but the 90% confidence range for each abundance parameter becomes larger. The best-fit parameters for the lower temperature component are $T:0.811^{+0.004}_{-0.004}$, Fe:$1.89^{+0.67}_{-0.57}$, O:$1.29^{+0.43}_{-0.65}$, while values for the higher temperature component are $T:1.445^{+0.05}_{-0.05}$, Fe:$1.20^{+0.47}_{-0.13}$, and O:$0.10^{+0.68}_{-0.10}$. These two sets of abundances do not show much difference within their 90% error bars. They are also consistent with the best-fit parameters when linking the abundances of the 2T models.

So in our sample, even for the region that has the most contamination from the cluster emission, linking the two abundance sets would not significantly affect the hot gas abundances, and it can better constrain the abundances by using fewer free parameters. Direct measurements of abundances for different thermal components are not good for low-quality spectra, and a deprojected method should be used.
3.6. Discussion on the Effect of RGS Background

For the 1′ RGS region, the local galaxy hot gas emission is always the dominant component. Thus it makes little difference, whether using an RGS background template or using one that contains the underlying cluster emission but free of local galaxy hot gas emission.

For the 3′5 RGS region, however, the background could have some effect, particularly for those with the brightest surrounding group/cluster emission. In our sample, fainter systems such as NGC4552, NGC720, NGC3923, and IC1459 have too poor RGS spectra to constrain their abundances, even for the 3′5 region. For brighter systems, however, it is very hard to get a local background that contains only the underlying cluster emission but free of galaxy hot gas emission. This is because their extended hot gas emission regions are usually larger than 5′ in diameter (checked with XMM-Newton MOS images), but the maximum field of view of RGS in the cross-dispersion direction is only 5′. The only way to obtain this local background is using separate observation pointing off the target and on the surrounding cluster emission, which is not available for all the galaxies in our sample. Thus using local contaminated background will overestimate the background, while using the RGS background template will underestimate the underlying cluster emission. This is the main error source when determining the abundances for the RGS 3′5 region, especially when the cluster hot gas component is comparable to the galaxy hot gas component, which can be seen by the luminosity ratio of the two thermal components obtained from the EPIC fits.

4. DISCUSSION AND CONCLUSIONS

There have been several previous determinations of metal abundances in early-type galaxies, some of which have obtained similar results (Buote 1999; Humphrey & Buote 2006; Athey 2007). One important difference in this work is that, for a sample of 10 galaxies, we include the RGS data, in which one can clearly see the strong individual lines from Fe, O, and other species, whereas they are thoroughly blended in the XMM-Newton EPIC and in the Chandra ACIS data. The identification of these ions not only gives us greater confidence in the abundance determinations, but they often provide very strict constraints on the maximum allowable abundances.

Of the models considered, the one yielding the lowest $\chi^2$ has a power-law to represent unresolved X-ray binaries, two thermal components with one set of variable abundances for
several individual elements (Fe, O, Mg, and Si), and Galactic absorption. This model led to formally acceptable or good fits for most of the galaxies except for NGC 4636 and IC 1459. In IC 1459, there is a strong central point source that compromised spectral fits to the fainter emission from the hot gas. In NGC 4636, a 2T model is inadequate, possibly due to the structure and events that have been described by others (Jones et al. 2002; O’Sullivan et al. 2005). Excluding NGC 4636 and IC 1459, the remaining eight galaxies have median abundances (25% and 75% quartile points are given in parenthesis for Fe and O) relative to solar values (Grevesse & Sauval 1998) of 0.86 (0.64–1.70), 0.44 (0.29–0.60), 0.81, and 0.79 for Fe, O, Mg, and Si. The abundances for Mg and Si are not distinguishable from that of Fe. Three of the galaxies have mildly supersolar Fe abundances, at about 1.5 of the solar value, and we note that the three all lie at the center of their group or cluster (NGC 1399 in the Fornax cluster; NGC 4472 is the dominant galaxy in the southern part of the Virgo cluster; NGC 4649 is the brightest member of a group on the outskirts of the Virgo cluster). While this may be a trend, we note that the galaxy NGC 5044 is also the brightest member of its group yet it has abundances that are close to the median values. The galaxy that has the lowest Fe abundance is NGC 4552, the optically least luminous galaxy of this group (and relatively lower in X-ray luminosity). As we will discuss in a future paper, relatively low abundances are found frequently in galaxies with lower X-ray and optical luminosities, as O’Sullivan & Ponman (2004) has found.

One of the issues that we examined was the difference between abundances when a second temperature component is added, sometimes referred to as the Fe-bias (Trinchieri et al. 1994; Buote 2000). The need for a second temperature component is probably due to a temperature gradient within the galaxy. Upon adding the second thermal component, the increase in Fe was about 30% (quartile values of 7%–38%) and for O it was 4% (quartile values of −3% to 16%). There are two galaxies in which there was essentially no change in the abundances, NGC 4649 and NGC 5044. From radial temperature profiles, NGC 4649 is known to be nearly isothermal (Athey 2007; Randall et al. 2006), which explains why adding a second temperature component was of little consequence. For NGC 5044, there was always a dominant thermal component and the luminosity from an additional thermal component was relatively small in the best-fit solutions.

Because NGC 4649 can be adequately fit with a single-temperature component and the temperatures do not change with an additional thermal component, we used this source to compare results between spectral fits for the three XMM-Newton instruments (EPIC MOS, EPIC PN, and the RGS) and for the Chandra ACIS-S. The derived temperatures were the same to about 1%, and there was excellent agreement between the three XMM-Newton instruments for the seven fitted abundances. The abundances derived from the Chandra ACIS-S were also in good agreement for Fe, Mg, Si, and Ni, but gave lower values than the XMM-Newton data for O, Ne, and S. Overall, Fe, Mg, and Si had similar abundances that were 1.5 times the solar value, S is consistent with that value, and Ni may be supersolar, although the constraints are poor. Oxygen, the lightest metal in the group, has an abundance of about 0.6 solar and Ne, the next lightest metal, has an abundance between O and Fe, at about 0.7 solar.

Metals in hot gas of early type galaxies are thought to come from stellar mass loss, along with SNe. Since SNe mainly contribute heavy elements to the hot gas, we should expect supersolar metallicity in hot gas. In Figure 10, we compared Fe abundances found in hot gas and those from optical studies (only six galaxies in our sample have optical Fe abundances reported in literature by Humphrey & Buote 2006 and Trager et al. 2000). Only three galaxies (NGC 1399, NGC 4472, and NGC 4649) show modest supersolar Fe abundances in hot gas, while the other three (NGC 720, NGC 3923, and NGC 4552) are slightly subsolar. SNe were expected to enhance the gas phase abundances considerably, which is not observed in our sample.

Element abundance ratios can constrain the relative contributions of SN Ia or SN II to the metallicity. In our sample, the median abundance ratios are 0.51, 0.94, and 0.92 for O/Fe, Mg/Fe, and Si/Fe, respectively. These give us the fraction of contribution from Type Ia SNe of 64%–85% when using linear combination of SN Ia yields and SN II yields from Nomoto et al. (1997), which indicates that most of the metal enrichment is from Type Ia SNe.

One important result is the low O/Mg ratio, which we derived to be a value of 0.54, which is also confirmed to be low by Athey (2007) and Humphrey & Buote (2006). Theoretically, one should expect similar abundance of O to Mg, since SN Ia produces a similar amount of O and Mg while SN II produce slightly more O than Mg. Chemical studies of Galactic bulge stars, which are also old star population similar to stars in elliptical galaxies but are mainly enriched by Type II SNe, also show a similar low O/Mg ratio at high Fe abundance (Fulbright et al. 2004; Minnit & Zoccali 2008). This low O abundance in the hot gas as well as in bulge stars might be the evidence of an overestimate of O yield by SN II models, which do not consider significant mass loss at the late stage of massive progenitor stars. Since strong stellar wind from massive progenitor stars might reduce or even completely blow away their He-rich envelopes while still preserve the C–O layers, thus after Type II SN explosions, only α elements (like Mg) by Carbon burning are synthesized while elements like O through hydrostatic He burning are greatly reduced (Fulbright et al. 2004). The observed Mg-rich SN remnant of N49B without accompanying Ne or O can support this argument (Park et al. 2003).
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