MOBSTER – V. Discovery of a magnetic companion star to the magnetic β Cep pulsator HD 156424

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ABSTRACT

HD 156424 (B2 V) is a little-studied magnetic hot star in the Sco OB4 association, previously noted to display both high-frequency radial velocity (RV) variability and magnetospheric Hα emission. We have analysed the Transiting Exoplanet Survey Satellite (TESS) light curve, and find that it is a β Cep pulsator with 11 detectable frequencies, 4 of which are independent p-modes. The strongest frequency is also detectable in RVs from ground-based high-resolution spectroscopy. RVs also show a long-term variation, suggestive of orbital motion with a period of ~years; significant differences in the frequencies determined from TESS and RV data sets are consistent with a light-time effect from orbital motion. Close examination of the star’s spectrum reveals the presence of a spectroscopic companion, however, as its RV is not variable it cannot be responsible for the orbital motion and we therefore infer that the system is a hierarchical triple with a so-far undetected third star. Reanalysis of least-squares deconvolution profiles from ESPaDOnS and HARPSpol spectropolarimetry reveals the surprising presence of a strong magnetic field in the companion star, with ⟨Bz⟩ about +1.5 kG as compared to ⟨Bz⟩ ~−0.8 kG for the primary. HD 156424 is thus the second hot binary with two magnetic stars. We are unable to identify a rotational period for HD 156424A. The magnetospheric Hα emission appears to originate around HD 156424B. Using Hα, as well as other variable spectral lines, we determine a period of about 0.52 d, making HD 156424B one of the most rapidly rotating magnetic hot stars.

Key words: binaries: spectroscopic – stars: early-type – stars: individual: HD 156424 – stars: magnetic field – stars: oscillations.

1 INTRODUCTION

Approximately 10 per cent of early-type stars are magnetic (e.g. Grunhut et al. 2017; Sikora et al. 2019a). In contrast to the magnetic fields of stars with convective envelopes, which are sustained by contemporaneous rotational–convective dynamos (e.g. Donati et al. 2008; Folsom et al. 2016, 2018), no dynamo mechanism has been confirmed to be sustainable in the radiative envelopes of hot stars. This has led to the suggestion that massive star magnetic fields are ‘fossils’, remnants of a previous era in the star’s formation that retain their stability in highly conductive radiative envelopes (e.g. Braithwaite & Spruit 2004; Braithwaite 2009; Duez, Braithwaite & Mathis 2010; Neiner et al. 2015). A fossil origin is consistent with the observed properties of hot star magnetic fields: they are topologically simple (typically ‘distorted dipoles’; Kochukhov, Shultz & Neiner 2019); stable over at least decades (e.g. Shultz et al. 2018b); their unsigned magnetic flux is either conserved (Sikora et al. 2019b) or slowly decays (Landstreet et al. 2007, 2008; Fossati et al. 2016; Shultz et al. 2019b) over evolutionary time-scales; and there is no clear correlation between the surface magnetic field strength and physical properties such as rotation (e.g. Shultz et al. 2019b; Sikora et al. 2019b), in contrast to what is unambiguously observed for stars with convective envelopes (e.g. Folsom et al. 2016, 2018).

The external properties of magnetic hot stars are relatively well understood, but little is known about the internal magnetic configurations of stars with fossil fields, resulting in great uncertainty regarding evolutionary models incorporating fossil magnetism (Keszthelyi et al. 2019, 2020; Takahashi 2020). There is some evidence that extremely strong magnetic fields can suppress convection in opacity-bump convection zones, and thereby reduce or eliminate macroturbulence (Sundqvist et al. 2013). Asteroseismological analysis has suggested that core overshooting may be suppressed by fossil magnetic fields (Briquet et al. 2012). Since the best method of probing stellar interiors is via asteroseismology, identification of appropriate magneto-asteroseismic targets is a high priority (e.g. Bowman et al. 2018; Buysschaert et al. 2018, 2019). Efforts to expand the sample of asteroseismic targets on the upper main sequence are an important part of this, since about 10 per cent of them can be expected to be magnetic (e.g. Burssens et al. 2019, 2020; Pedersen et al. 2019; Labadie-Bartz et al. 2020).

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HD 156424 is a B2 V star in the Sco OB4 association. Alecian et al. (2014) detected the star’s magnetic field, finding a longitudinal magnetic field of about $-500$ G. Further observations were analysed by Shultz et al. (2018b), who confirmed the low level of variation in the magnetic field and reported a tentative periodicity of about 2.8 d. Radial velocity (RV) variation of a few km s$^{-1}$ was reported by Alecian et al. (2014).

Alecian et al. (2014) also reported the presence of H$\alpha$ emission consistent with an origin in a centrifugal magnetosphere (CM; e.g. Landstreet & Borra 1978; Townsend & Owocki 2005; Petit et al. 2013). In most cases, CM-type H$\alpha$ emission occurs above the Kepler corotation radius $R_K$, which for the rapidly rotating, strongly magnetized stars that host such emission is generally at a distance of a few stellar radii (e.g. Shultz et al. 2019b), or equivalently a few times $v\sin i$. HD 156424’s H$\alpha$ emission is anomalous in this regard because it peaks at about 20–30 times $v\sin i$, apparently much larger than the star’s Alfvén radius (Shultz et al. 2019b), i.e. at a greater distance than the maximum extent of magnetic confinement (ud-Doula & Owocki 2002). HD 156424’s magnetosphere has also been detected in gyrosynchrotron emission (Kurapati et al. 2017), although not yet in X-rays (Naz´ee ta l. 2014).

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The rapid RV variation opens the possibility that HD 156424 may be one of the rare class of pulsating magnetic stars. Determining if this is the case, and if so providing an initial characterization of its pulsation properties, provided the initial motivation for our analysis of the recently obtained Transiting Exoplanet Survey Satellite (TESS; Ricker et al. 2015) light curve together with the existing ground-based spectroscopic data set. Further motivation was provided by the anomalous nature of the H$\alpha$ emission noted by Shultz et al. (2020), and the ambiguous rotation period inferred from magnetic data by Shultz et al. (2018b). A description of the available data sets is provided in Section 2, while the frequency analyses of the light curve and RVs are, respectively, described in Sections 3 and 4. The RV analysis provided unexpected evidence that the star has a companion, in consequence of which we analyse the spectra looking for evidence of binarity in Section 5, following which we determine the atmospheric properties of the components. A reanalysis of the magnetic and H$\alpha$ data in the light of multiplicity is provided in Sections 6 and 7, with the surprising results that both stars are magnetic, and that the H$\alpha$ emission almost certainly originates around the secondary. Magnetic models are inferred in Section 8. The conclusions are summarized in Section 9.

2 OBSERVATIONS

2.1 ESPaDOnS spectropolarimetry

ESPaDOnS is a fibre-fed echelle spectropolarimeter mounted at the Canada–France–Hawaii Telescope (CFHT). It has a spectral resolution $\lambda/\Delta\lambda \sim 65,000$, and a spectral range from 3700 to 10,500 Å over 40 spectral orders. Each observation consists of four polariometric sub-exposures, between which the orientation of the instrument’s Fresnel rhombs are changed, yielding four intensity (Stokes I) spectra, 1 circularly polarized (Stokes V) spectrum, and two null polarization (N) spectra, the latter obtained in such a way as to cancel out the intrinsic polarization of the source. Wade et al. (2016) describe the reduction and analysis of ESPaDOnS data in detail. Nine Stokes V observations were acquired between 04/2014 and 06/2014 by a P.I. program.\footnote{Program Code CFHT 14AC010.} A uniform sub-exposure time of 450 s was used for all observations. The median peak signal-to-noise ratio (S/N) per spectral 1.8 km s$^{-1}$ pixel is 369.

2.2 HARPSpol spectropolarimetry

HARPSpol is a high-resolution ($\lambda/\Delta\lambda \sim 110,000$) echelle spectropolarimeter with a spectral range covering 3780–6910 Å, with a gap between 5240 and 5360 Å, across 71 spectral orders. It is installed at the 3.6-m telescope at the European Southern Observatory (ESO) La Silla facility. As with ESPaDOnS, each spectropolarimetric sequence consists of four polarized sub-exposures, which are combined to yield the Stokes $V$ spectrum as well as a diagnostic null $N$. The sub-exposure time was 900 s. Three observations were acquired in 2012 by the Magnetism in Massive Stars (MiMeS) ESO Large Program. The acquisition, reduction, analysis, and characteristics of these data was described by Alecian et al. (2014).

2.3 FEROS spectroscopy

FEROS is a high-dispersion echelle spectrograph, with $\lambda/\Delta\lambda \sim 48,000$ and a spectral range of 3750–8900 Å (Kaufer & Pasquini 1998). It is mounted at the 2.2-m La Silla MPG telescope. We acquired 11 spectra between 2015 June and 2015 July, with an exposure time of 1400 s. The data were reduced using the standard FEROS Data Reduction System MIDAS scripts.\footnote{Available at https://www.eso.org/sci/facilities/lasilla/instruments/feros/tools/DRS.html.} The median peak S/N per 1.4 km s$^{-1}$ spectral pixel is 261.

2.4 TESS photometry

TESS is a space telescope obtaining high-precision ($\mu$mag) photometry (Ricker et al. 2015). Its initial mission will last 2 yr, during which it will observe 85 percent of the sky in overlapping sectors of $96 \times 24^\circ$. Each sector is observed for about 27 d. Data for high-priority targets is downloaded with a 2-min cadence. The instrument obtains data over a broad bandpass (6000–10,000 Å), with large (21 $\times$ 21 arcsec) pixels.

HD 156424 was observed by TESS in Sector 12 using 2-min cadence, with a total of 15 784 individual observations. We obtained the light curve from the Mikulski Archive for Space Telescopes (MAST), selecting the pDCSAP flux as the light curve with the best apparent detrending. No additional detrending was required, since we are not interested in long-term trends. While the star lies in a relatively crowded field and there are certainly other stars contaminating the TESS light curve, it is the brightest star within about 4 arcmin.

3 PHOTOMETRIC ANALYSIS

The TESS light curve is shown in the right-hand panel of Fig. 1 and evidences clear, multiperiodic variability. Analysis of the light curve with PERIOD04 (Lenz & Breger 2005) reveals several significant frequencies. These are listed in Table 1. The four strongest frequencies have amplitudes of about 2–3 mmag, and are indicated in the top left-hand panel of Fig. 1. After pre-whitening with these frequencies, seven more frequencies with amplitudes of about 0.1 mmag are detected; these are shown in the bottom left-hand panel of Fig. 1. Significance was determined according to the usual criterion of a S/N of at least 4 (Breger et al. 1993; Kuschnig et al. 1997). This noise floor is shown in Fig. 1, where it was determined by fitting...
Table 1. Frequencies from the TESS light curve and radial velocity measurements, with uncertainties in the last digit given in parantheses. TESS amplitudes are in mmag, with an uncertainty of about 0.01 mmag; RV amplitudes are in km s\(^{-1}\), with an uncertainty of about 0.1 km s\(^{-1}\). The third column gives the signal-to-noise ratio (S/N) of the frequency. The final column gives the identification.

| Label | Frequency (c d\(^{-1}\)) | Amplitude (mmag) | S/N | ID |
|-------|--------------------------|------------------|-----|----|
| TESS  |                          |                  |     |    |
| \(f_1\) | 11.20672(7) | 3.27             | 254 |   |
| \(f_2\) | 0.71699(8)  | 3.10             | 24  |   |
| \(f_3\) | 13.7753(1)  | 2.11             | 177 |   |
| \(f_4\) | 10.8634(1)  | 1.75             | 130 |   |
| \(f_5\) | 1.4291(7)   | 0.35             | 4   | \(2f_2\) |
| \(f_6\) | 11.5015(8)  | 0.30             | 22  | \(2f_1 - f_4\) |
| \(f_7\) | 22.363(1)   | 0.21             | 18  | \(2f_1\) |
| \(f_8\) | 24.639(3)   | 0.06             | 6   | \(f_3 + f_4\) |
| \(f_9\) | 16.187(4)   | 0.06             | 5   |   |
| \(f_{10}\) | 24.991(4)  | 0.05             | 5   | \(f_1 + f_3\) |
| \(f_{11}\) | 22.406(4)  | 0.05             | 5   | \(2f_1\) |
| FEROS |                          |                  |     |    |
| \(f_1\) | 11.213(7)   | 3.3              | 12  |   |
| ESPaDOnS |          |                  |     |    |
| \(f_1\) | 11.20672(2) | 4.0              | 28  |   |
| HARPSpol |           |                  |     |    |
| \(f_1\) | 11.17(1)    | 3.4              | 27  |   |
| Combined spectroscopy | |                  |     |    |
| \(f_1\) | 11.20691(1) | 3.7              | 33  |   |
| \(f_2\) | 12.22456(5) | 0.8              | 7   |   |

The majority of the frequencies are above 10 d\(^{-1}\), and are consistent with \(p\)-mode pulsations. Three of the \(p\)-modes \((f_1, f_2, \text{and } f_6)\), including the strongest frequency, constitute a triplet centered on \(f_1\), with a separation of about 0.3 d\(^{-1}\). With the exception of \(f_9\), the remaining high-frequency terms all appear to be either harmonics or linear combinations of the strongest frequencies, as determined using the Rayleigh criterion with a full width at half-maximum of about 0.06 c/d.

In addition to the \(p\)-modes, there are two low-frequency terms in the frequency spectrum, the lowest of which, \(f_2\), has the second-highest amplitude of all detected frequencies. The weaker signal at \(f_4\) is very close to the first harmonic of the stronger low-frequency term \(f_2\), suggesting that this might be due to rotational modulation. The period corresponding to \(f_2\), about 1.4 d, is about half the period reported by Shultz et al. (2018b). This could indicate that \(f_2\) is in fact the first harmonic of \(f_{\text{rot}}\); however, there is no statistically significant peak at the corresponding frequency. Therefore, if \(f_2 = 2f_{\text{rot}}\), the rotationally modulated light curve would need to be an almost perfect double-wave variation. It is interesting to note that there are two sets of frequencies \((f_1 - f_2\) and \(f_2 - f_0)\) that are separated by about 0.35 d\(^{-1}\), which is very close to the presumptive \(f_{\text{rot}} = f_2/2\). The elements of the triplet centred on \(f_1\) are also separated by about 0.3 d\(^{-1}\). These frequency groups may therefore be a consequence of frequency splitting, supporting the interpretation of rotational modulation for \(f_2\). On the other hand, this could be coincidental, and the low-frequency terms might be due to \(g\)-mode pulsations.
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4 RADIAL VELOCITIES

Radial velocities (RVs) were measured from individual unpolarized spectra, yielding 36 RV measurements from the ESPaDOnS data, 12 from the FEROS data, and 12 from the HARPSpol data. RVs were measured via the centre-of-gravity of the Si III λλ 4553, 4568, and 4575 lines, as well as the O II λλ 4415 line. As explained below in Section 5, these lines are dominated by the spectrum of the primary. All yielded similar results within uncertainties. The weighted mean across all 4 lines was then taken so as to increase the S/N. RVs are tabulated in Table A1.

RVs are shown as a function of time in the top panel of Fig. 2. It is immediately apparent that there is a systematic difference between the HARPSpol RVs and the RVs measured from the ESPaDOnS and FEROS data sets, with a difference of about 7 km s\(^{-1}\). The most natural explanation for this is orbital motion due to the influence of a companion star.

Fourier analysis of the ESPaDOnS RVs using PERIOD04 yielded a single significant frequency, 11.207(2) d\(^{-1}\), with an S/N of 28. After pre-whitening with this frequency, the next highest peak has a S/N below the significance threshold of 4. The FEROS measurements yield 11.213(7) d\(^{-1}\), with a S/N of 12. The HARPSpol data yields 11.17(1) d\(^{-1}\), with a S/N of 27. Combining the three data sets (after removing the mean RV of each data set so as to correct for the systematic differences) yields 11.20691(1) d\(^{-1}\). A second frequency is found at 12.22456(5) d\(^{-1}\), although the absence of this frequency in the much more precise photometric data set suggests it may be spurious (it is also worth noting that the amplitude is similar to the RV uncertainty). RVs are shown phased with \(f_1\) (as determined from RVs) in the bottom panel of Fig. 2.

In \(\beta\) Cep pulsators, photometric variations occur primarily due to changes in \(T_{\text{eff}}\) with pulsation phase, which can sometimes also be detected as changes in line strength. Fig. 3 shows a dynamic spectrum for Si III λλ 4553, with individual spectra moved to the laboratory frame and folded with \(f_1\). With RV variation removed, there is no apparent line profile variability down to about 1 percent of the continuum. This is probably consistent with the very low photometric and RV amplitude of \(f_1\) since intrinsic variation in line profile strength is due to the change in surface temperature, which is minimal in this case.

5 MULTIPLICITY

A companion star is known to be about 0.35 arcmin away (Hartkopf et al. 1993; Tokovinin, Mason & Hartkopf 2010). Sco OB4 is estimated to be at a distance of 1300\(^{+500}_{-200}\) pc (Kharchenko et al. 2005), so the projected separation of the stars is 455\(^{+175}_{-70}\) au. The magnitude...
was performed using a grid-based search, covering 20kK from the BSTAR2006 library (Lanz & Hubeny 2007). The fitting spectral lines are shown in the top panels of Fig. 4. Since the for which different ionizations are available in the spectrum. These examined a selection of metallic lines, including especially lines to Si III, this pattern is clearly apparent. Lines with lower ionizations have much more extended wings than lines of higher ionizations. Comparing O I to O II, Fe II to Fe III, and Si II these fit due to the presence of circumstellar emission. Bottom panels show residual flux; note the presence of emission in all three lines.

Figure 4. Top: Synthetic TLUSTY spectrum fits to metallic lines in the mean ESPaDOnS spectrum. Primary in dashed blue, secondary in dot–dashed purple, and combined in solid red. Primary and secondary spectra are intrinsic, i.e. they have not been scaled by luminosity. Note the difference in line profile shape between low ionization lines (O I, Si II, Fe II) with stronger contributions from the secondary, and high ionization lines (O II, Si III, Fe III) with stronger contributions from the primary. Bottom: Synthetic TLUSTY spectrum fits to Balmer lines in the mean ESPaDOnS spectrum. The vertical dotted lines delineate regions excluded from the fit due to the presence of circumstellar emission. Bottom panels show residual flux; note the presence of emission in all three lines.

difference is estimated at 2.3 mag in the y band (Tokovinin et al. 2010), or about a factor of 8 in luminosity (implying the companion star should have log L ∼ 2.8). In this case it should have a mass of about 5 or 6 M☉, for a mass ratio of ∼ 3. From Kepler’s third law, the orbital period would then need to be on the order of 200 yr. However, speckle observations indicate the companion’s position angle changed by 30° between 1990 and 2008 (Hartkopf et al. 1993; Tokovinin et al. 2010); this would suggest an orbital period closer to 200 yr. In either case, this is much too long to be consistent with the observed long-term RV variation.

Alecian et al. (2014) indicated that they found no indication of the companion star in the spectrum. However, given the companion star’s angular distance from the primary, it was inside the 2 arcmin FEROS aperture, the 1 arcmin HARPSpol aperture, and the 1.8 arcmin ESPaDOnS pinhole. To conduct a more detailed investigation, we created a mean spectrum from all available ESPaDOnS spectra, achieving a peak S/N per 1.8 km s⁻¹ pixel of about 800. We then examined a selection of metallic lines, including especially lines for which different ionizations are available in the spectrum. These spectral lines are shown in the top panels of Fig. 4. Since the companion star is estimated to be about 2.3 mag dimmer than the primary, it should have a significantly lower T eff, and should therefore contribute different amounts to the flux of lines with different ionizations. Comparing O I to O II, Fe II to Fe III, and Si II to Si III, this pattern is clearly apparent. Lines with lower ionizations have much more extended wings than lines of higher ionizations.

The top panels of Fig. 4 show a fit to a selection of metallic lines in the mean ESPaDOnS spectrum using synthetic TLUSTY spectra from the BSTAR2006 library (Lanz & Hubeny 2007). The fitting was performed using a grid-based search, covering 20kK < T eff, A < 24 kK for the primary, 15kK < T eff, B < 20 kK, 3.5 < log g < 4.5 for both stars, 1 < vsin i A < 10 km s⁻¹, and 10 < vsin i B < 100 km s⁻¹. The radius ratio R A/R B was allowed to vary as a free parameter, and was calculated using the continuum fluxes from the TLUSTY spectra. RVs were fixed to 8 and 5 km s⁻¹ for A and B, respectively. The best fit was obtained for log g = 4.25 for both stars, T eff, A = 23 ± 1 kK, T eff, B = 16 ± 1 kK, vsin i A = 4.4 ± 1.5 km s⁻¹, and vsin i B = 25 ± 2 km s⁻¹. The best-fitting value of R A/R B = 1.1 certainly overestimates the radius of secondary, which may be a consequence of chemical peculiarities in one or both stars. This results in HD 156424A being about twice as bright as HD 156424B at visible wavelengths, whereas the magnitude difference estimated by Tokovinin et al. (2010) implies the primary should be about eight times brighter than the secondary; however, Tokovinin et al. noted that the magnitude difference is probably overestimated due to the noisy data.

The bottom panels of Fig. 4 show fits to the H Balmer lines H γ, H β, and H α. In this case the T eff of the two components were fixed to their best-fit values. Due to the presence of obvious circumstellar emission in H α, the regions with emission were excluded from the fit; the corresponding regions in H γ and H β were also excluded. The best fit is obtained for log g A = 4.0 and log g B = 3.75. This is a curious result since, assuming the two stars to be coeval, the surface gravity of HD 156424B should be at least as high, or higher, than that of the more massive primary. However, if the grid is restricted to force log g A < log g B, the difference in fit quality is negligible, and it seems likely that, given the small RV separation of the components, their respective surface gravities cannot be confidently disentangled.

RV measurements of the secondary were conducted using the parametrized line profile fitting program described by Grunhut et al. 

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An eccentric orbit would indicate a lower mass for the companion.

be appropriate for a young star with its eccentricity is 0 the companion must have a mass of at least 1 M⊙ the maximum of an approximately sinusoidal variation. We further change through this time span, suggesting the RV curve is near between the ESPaDOnS and FEROS data indicates a low rate of mean RV had not yet returned to the HARPSpol value at the time by Hartkopf et al. (1993) via speckle interferometry.

seems likely that the primary’s long-term RV variation is caused by an undetected third star, and that the system is a hierarchical triple, with HD 156424A and the undetected companion in an orbit with a period of ~years, and HD 156424B orbiting the inner pair with an orbital period of centuries (as inferred from visual data). In this scenario, HD 156424B is the distant companion previously identified by Hartkopf et al. (1993) via speckle interferometry.

A lower limit can be placed on the mass of the undetected companion if we assume that the observed RV variation samples about half of an orbital period (in which case the orbital period would be about 5 years), and the semi-amplitude of the orbital RV variation is therefore about 4 km s⁻¹. Half a period is assumed because the mean RV had not yet returned to the HARPSpol value at the time of the FEROS observations; at the same time, the lack of variation between the ESPaDOnS and FEROS data indicates a low rate of change through this time span, suggesting the RV curve is near the maximum of an approximately sinusoidal variation. We further assume that HD 156424A has a mass of about 8 M⊙, as would be appropriate for a young star with its Teff. By iteratively solving for the mass of the unseen companion using the mass function of the system (e.g. Batten 1973; see also equations A-1 and A-2 in Rivinius et al. 2020), if the orbital inclination is 90° and the eccentricity is 0 the companion must have a mass of at least 1 M⊙. An eccentric orbit would indicate a lower mass for the companion. Since a 1 M⊙ star would have a bolometric luminosity of less than a thousandth that of HD 156424A, it is entirely possible for such a star to remain undetected. An approximate upper limit on the mass of the companion can be placed under the assumption that it would be seen if its mass were comparable to that inferred for HD 156424B from its Teff at the ZAMS, about 6 M⊙; in this case, the orbital inclination must be greater than 13°.

There is an apparent change in f₁ between the combined RVs and the TESS light curve. This is significant at almost the 3σ level using the larger TESS uncertainty. The change amounts to ΔP ~ 1.5 × 10⁻⁶ d. If this is due to the light-time effect, there should be a corresponding change in RV of ΔRV = cΔP/P ~ 5 km s⁻¹ (Pigulska 1992). Referring to Fig. 2, this is about the change in RV that is observed between the acquisition of the HARPSpol and ESPaDOnS/FEROS data sets. The apparently much larger difference in f₁ as determined from the HARPSpol data set and others is also formally significant at the 3σ level using the larger HARPSpol uncertainty, but would require ΔRV ~ 1000 km s⁻¹ to be explained via the light-time effect. This change, if real, therefore cannot be explained by the light-time effect; however, as it is based on a relatively small number of measurements that are phased equally well using a period closer to the mean value, this large difference in frequencies is probably spurious.

### 6 MAGNETOMETRY

In order to measure the magnetic field with the maximum possible precision, mean line profiles were extracted from the spectropolarimetric data using least-squares deconvolution (LSD; Donati et al. 1997), for which we employed the ILSD package (Kochukhov, Makaganiuk & Piskunov 2010). Line masks were obtained from the Vienna atomic line data base (Piskunov et al. 1995; Ryabchikova et al. 1997; Kupka et al. 1999, 2000; Ryabchikova et al. 2015) using ‘extract stellar’ requests. The first such line mask was obtained for the stellar parameters of HD 156424 (Teff = 20 K, log g = 4) reported by Shultz et al. (2019a), with a normalized line depth threshold of 0.1. The line mask was then prepared in the usual way: cleaned of contaminating H lines, He lines with broad wings, telluric lines, and interstellar lines (including 2 Diffuse Interstellar Bands at 5780 and 6604 Å), and the line strengths adjusted to match the observed line depths (e.g. Shultz et al. 2018b).

The presence of a magnetic field was evaluated using False Alarm Probabilities (Donati, Semel & Rees 1992; Donati et al. 1997). All Stokes V profiles yield ‘Definite Detections’, i.e. FAP < 10⁻⁵. However, 8/12 N profiles also yield DDS, the remainder being non-detections (NDS; FAP > 10⁻⁵). These detections in N are due to the high-frequency RV variation from β Cep pulsations. Subexposure times for the ESPaDOnS data are about 6 per cent of f₁, and total spectropolarimetric sequence times are 23 per cent of f₁. HARPSpol sequences, all three of which show signatures in N, span 47 per cent of a pulsation cycle. This phenomenon has been reported in other stars with rapid RV variation, e.g. the β Cep stars HD 96446 and η¹ CMa (Neiner et al. 2012b; Shultz et al. 2017), and the short-period binary HD 156324 (Shultz et al. 2018a). While Stokes V is undoubtedly also affected, the total strength of the line-of-sight integrated magnetic field (B₁) (defined by Mathys 1989) should not be affected, as determined via correction of supexposures for RV variation (Neiner et al. 2012b) and via direct modelling of Stokes I, Stokes V, and N (Shultz et al. 2017).

Close examination of the Stokes V profiles, however, reveals a curious feature: lobes of net polarization extending outside the line profile (Fig. 6, top). To investigate the effect of these anomalous polarization lobes, (B₁) was measured from the mean line profile using progressively wider integration ranges. In general, (B₁) should change up to the edge of the line width, following which it should
Figure 6. Mean LSD profiles extracted with various line masks. The grey-shaded regions indicate the mean uncertainties in $N$ and Stokes $V$.

stabilize while the uncertainty continues to grow (e.g. Neiner et al. 2012a). As can be seen in Fig. 7, for the LSD profiles extracted using the 20 kK mask $\langle B_z \rangle$ continues to change outside of the error bars well after this point, and indeed reverses sign around 40 km s$^{-1}$. This behaviour is quite anomalous and cannot be explained as a consequence of pulsational influence on the line profile, which should only be able to shift the Stokes $V$ profile by a few km s$^{-1}$.

The lobes in Stokes $V$ correspond to the very broad wings of Stokes $I$. Since He lines were excluded from the line mask, the presence of these wings cannot be explained as due to Stark broadening. The ‘wings’ may instead be explained by the contribution of HD 156424B to the spectrum.

The strange behaviour of $\langle B_z \rangle$, and the correspondence between the ‘wings’ in Stokes $I$ and the lobes in Stokes $V$, suggest that HD 156424B may possess its own magnetic field and therefore contaminates both Stokes $I$ and $V$. Given its likely $T_{\text{eff}}$ of about 16 kK, if the star is indeed magnetic it is almost certainly a He-weak Bp star. A line list was downloaded from VALD using enhanced abundances $\{\text{Si} = -4, \text{Cl} = -5, \text{Ti} = -5.5, \text{Cr} = -5, \text{Fe} = -3.8, \text{Ni} = -5.3, \text{Ba} = -7.8, \text{Ce} = -6, \text{Pr} = -8.2, \text{Nd} = -7.4\}$, chosen using the mean surface abundances for HR 2949 (Shultz et al. 2015a), which has a similar $T_{\text{eff}}$. To emphasize lines formed at lower $T_{\text{eff}}$ and therefore hopefully help to separate the signals, a $T_{\text{eff}}$ of 15 kK was used. The mask was then cleaned in the usual fashion to remove contamination from H Balmer, He, interstellar, and telluric lines. The mask was further cleaned so as to include only low-ionization (mostly Feii) lines, i.e. spectral lines in which the contribution from HD 156424B is dominant (see Fig. 4). Finally, a complementary mask was obtained using a solar abundance 26 kK line mask template, which was cleaned to remove any lines appearing in the 15 kK mask; in this case a higher $T_{\text{eff}}$ than that inferred from modelling of HD 156424A was chosen so as to emphasize lines appearing at higher ionizations that can be expected to minimize the contribution of the cooler star. LSD profiles were then extracted simultaneously using these two masks, a unique capability of the ILSD package (Kochukhov et al. 2010).

$\langle B_z \rangle$ measurements for the two sets of LSD profiles are listed in Table 2. $\langle B_z \rangle$ measurements are not affected by continuum dilution in binary star spectra, since they are are normalized using the EW of the Stokes $I$ profile (Mathys 1989); so long as Stokes $I$ is not contaminated by the other star, the measurement should therefore

Figure 7. Change in $\langle B_z \rangle$ determined from mean LSD profiles as a function of the integration range width. The vertical dashed blue and dot–dashed red lines indicate the line widths of the A and B components.
Table 2. \((B_\ast)\) and \((N_\ast)\) measurements for the two stellar components. The first column indicates the instrument with which the measurement was obtained: H(ARPSpol) or E(SPaDOnS). ‘DF’ refers to ‘detection flag’: either definite detection (DD), marginal detection (MD), or non-detection (ND).

| Inst. | HJD – 2456000 | Date       | \((B_\ast)\) (G) | DF\(\ast\) | \((N_\ast)\) (G) | DF\(\ast\) | \((B_\ast)\) (G) | DF\(\ast\) | \((N_\ast)\) (G) | DF\(\ast\) |
|-------|---------------|------------|------------------|-----------|------------------|-----------|------------------|-----------|------------------|-----------|
| H     | 126.66421     | 18/07/2012 | \(-849 \pm 26\) | DD        | \(-7 \pm 26\)   | DD        | 2154 \pm 274    | DD        | \(-415 \pm 270\) | ND        |
| H     | 126.70743     | 18/07/2012 | \(-901 \pm 38\) | DD        | \(-48 \pm 37\)  | DD        | 2586 \pm 391    | DD        | \(-114 \pm 385\) | ND        |
| H     | 127.77721     | 19/07/2012 | \(-686 \pm 48\) | DD        | \(-4 \pm 48\)   | DD        | 3005 \pm 392    | DD        | \(-403 \pm 385\) | ND        |
| E     | 758.03881     | 10/04/2014 | \(-846 \pm 54\) | DD        | 4 \pm 53         | ND        | 1645 \pm 418    | ND        | 249 \pm 416      | ND        |
| E     | 761.95376     | 14/04/2014 | \(-801 \pm 79\) | DD        | 43 \pm 79        | ND        | 1952 \pm 528    | ND        | \(-415 \pm 525\) | ND        |
| E     | 761.97745     | 14/04/2014 | \(-742 \pm 74\) | DD        | 0 \pm 74         | ND        | 3053 \pm 541    | ND        | \(-978 \pm 534\) | ND        |
| E     | 814.96460     | 06/06/2014 | \(-867 \pm 43\) | DD        | 10 \pm 43        | ND        | 1963 \pm 373    | ND        | \(-43 \pm 370\)  | ND        |
| E     | 814.98761     | 06/06/2014 | \(-821 \pm 37\) | DD        | 45 \pm 37        | DD        | 2281 \pm 337    | DD        | \(-326 \pm 333\) | ND        |
| E     | 821.94039     | 13/06/2014 | \(-818 \pm 36\) | DD        | \(-8 \pm 36\)    | DD        | 2777 \pm 335    | DD        | \(-79 \pm 328\)  | ND        |
| E     | 821.98111     | 13/06/2014 | \(-769 \pm 35\) | DD        | 26 \pm 35        | DD        | 1801 \pm 315    | DD        | \(-2 \pm 312\)   | ND        |
| E     | 824.89544     | 16/06/2014 | \(-859 \pm 33\) | DD        | 7 \pm 33         | DD        | 1234 \pm 266    | DD        | \(-128 \pm 265\) | ND        |
| E     | 824.91836     | 16/06/2014 | \(-853 \pm 34\) | DD        | \(-3 \pm 34\)    | DD        | 1533 \pm 393    | DD        | \(-280 \pm 379\) | ND        |

\(B_\ast\) measurements show very little variation for either component. For HD 156424A, the highest peak in the \((B_\ast)\) periodogram is at 1.539(2) d, corresponding to about 0.65 d\(^{-1}\); there is no peak in the TESS frequency spectrum corresponding to this frequency. However, the S/N of this peak is below 4 therefore it is not statistically significant. We are furthermore unable to confirm the 2.8 d period determined by Shultz et al. (2018b), and believe that this period was a spurious consequence of the unidentified contribution of the secondary. If we take the lowest frequency obtained from the TESS light curve, corresponding to a period of about 1.4 d, the \((B_\ast)\) measurements are not satisfactorily phased. We conclude that either a) the lowest frequency term in the TESS data set is not, in fact, the rotation period, or b) HD 156424B is still affecting HD 156424A’s Stokes V profile and the \((B_\ast)\) measurements are therefore not reliable. The latter conclusion seems more likely in light of the results of the integration range test (Fig. 7).

For HD 156424B, the strongest peak in the periodogram is at 0.8773(5) d, or about 1.14 d\(^{-1}\). Once again this does not correspond to any of the frequencies in the TESS light curve; however, the S/N of this frequency is below 4, therefore it is probably spurious. Similar to the case of the primary, the lowest frequencies in the TESS light curve do not satisfactorily phase \((B_\ast)\), and our conclusions for this star are identical. We conclude that \(P_{\text{rot}}\) cannot be determined using \((B_\ast)\) measurements alone for either star.

7 H\(\alpha\) EMISSION AND ROTATION

As can be seen from Fig. 9, and as was reported by Alecian et al. (2014), HD156424 displays H\(\alpha\) emission characteristic of a CM (Petit et al. 2013): two emission bumps at high velocities, presumably originating from two magnetically confined clouds within the warped disc of the CM (e.g. Landstreet & Borra 1978; Townsend & Owocki 2005). HD156424 is, however, an anomalous case. In general, the emission should be outside of the Kepler corotation radius \(R_\ast\); since the magnetically confined plasma is in strict corotation with the star, there is a linear mapping between line-of-sight velocity and projected distance. Assuming that the emission belongs to the primary, the emission peaks at about \(\pm 300\) km s\(^{-1}\) would then correspond to a projected distance of about 35 stellar radii. While this is outside the value of \(R_\ast = 4.2 R_\ast\) determined by Shultz et al. (2019b), it is also outside the Alfvén radius \(R_\text{A} = 21 R_\ast\), which is clearly impossible.
incorporating the flux of both components, assuming $T_{\text{RV}}$ variation of HD 156424, synthetic $H_{\alpha}$ available for HARPSpol. In order to minimize scatter due to the polynomial normalization process. Unnormalized spectra are not this was done to avoid warping of the line profiles due to the usual linear normalization performed on either side of the integration range; and FEROS data, un-normalized spectra were utilized, with a simple in an attempt to determine the rotational period. For the ESPaDOnS modulated purely by rotation. $H_{\alpha}$ as the host of the emission. Magnetic fields, making the secondary an inherently better candidate of HD 156424A’s Stokes $I$ profiles, suggesting that the line profile variability of HD 156424B is not a consequence of contamination by the Stokes $I$ profile of HD 156424A.

While these parameters change given the stronger $B_2$ and higher $T_{\text{eff}}$ inferred here for HD 156424, $R_2$ is very unlikely to become much larger since the higher $M$ will compensate for an increase in $B_2$ (a lower limit of 12 $R_\odot$ is inferred in Section 8).

If, on the other hand, the emission belongs to the previously unrecognized secondary, then the emission peaks occur at a somewhat more reasonable 11 $R_\odot$; while higher than the 3 or 4 $R_\odot$ at which CM emission peaks are generally observed, this is not unheard of, as in the case of CPD $-62\,2124$, for which Castro et al. (2017) found the emission peak to occur at about 8 $R_\odot$. HD 156424B is apparently rotating more rapidly than HD 156424A, and has a magnetic field at least twice as strong, with a minimum surface dipole strength of about 5 kG. Shultz et al. (2019b) showed that stars with CM-type $H_{\alpha}$ emission are exclusively rapid rotators with very strong magnetic fields, making the secondary an inherently better candidate as the host of the emission.

Assuming that the $H_{\alpha}$ emission is formed in the CM, it should be modulated purely by rotation. $H_{\alpha}$ equivalent widths were measured in an attempt to determine the rotational period. For the ESPaDOnS and FEROS data, un-normalized spectra were utilized, with a simple linear normalization performed on either side of the integration range; this was done to avoid warping of the line profiles due to the usual polynomial normalization process. Unnormalized spectra are not available for HARPSpol. In order to minimize scatter due to the RV variation of HD 156424, synthetic $H_{\alpha}$ spectra were calculated incorporating the flux of both components, assuming $T_{\text{effA}} = 23$ kK, $T_{\text{effB}} = 16$ kK, $\log g = 4.0$ in both cases, and shifting the individual spectra to their respective RVs. These fits are shown for maximum and minimum emission in the bottom panel of Fig. 9.

EWs were measured in the red and blue wings of the residual flux profiles, i.e. outside $\pm v \sin i$, with the outer limits defined by the maximum velocity of emission (i.e. about 500 km s$^{-1}$; the dotted green lines in Fig. 9). These emission EWs were analysed individually and combined; the periodogram for the total EWs is shown in Fig. 11. Maximum power is found at 0.52403(1) d for blue, 0.52681(2) d for red, and 0.524019(9) d for the total EWs.

Fig. 10 shows synthetic spectral TLUSTY fits to Fe II $\lambda\lambda$ 5169, Si II $\lambda\lambda$ 6347, Si II $\lambda\lambda$ 6371, and the O I $\lambda\lambda$ 7774 triplet, all lines in which the contribution of the secondary is relatively strong. The strength of the secondary’s contribution is variable: the left column shows the FEROS observation in which HD 156424B’s contribution is at a maximum, the right column a FEROS observation in which it is at a minimum. O and Si are more variable than Fe. This is probably due to the presence of chemical abundance patches on the secondary. Under the assumption that these line profile variations are due to spots, they should also show rotational modulation. We additionally measured the EWs of He I 6678 and Si II 6347, which are both relatively strong, and appear in all three spectroscopic data sets. Period analysis of the EWs measured from these lines finds maximum amplitude at 0.52399(1) d for He I 6678 (Fig. 11, bottom), and 0.53713(5) d for Si II 6347, essentially consistent with the results from $H_{\alpha}$.

Notably, the 1.39 d period obtained from the TESS photometry does not appear in any of the EW data sets (the red-dashed lines in Fig. 11). It therefore seems likely that the TESS period is not due to rotation, but may instead be due to gravity-mode pulsations in HD 156424A. Similarly, the 0.52 d period from the EWs does not appear in the TESS light curve. This may be a consequence of the light curve being dominated by the variability of HD 156424A, with the contribution of the dimmer HD 156424B only appearing in individual spectral lines in which its strong surface abundance patches lead to EW variations. Typical mCP stars have light-curve variations on the order of 1–10 mmag (e.g. David-Uraz et al. 2019; Sikora et al. 2019c), or about 0.1–1 percent of flux. Assuming HD 156424A is

![Fig. 8. Individual LSD profiles for HD 156424A (left) and HD 156424B (right). Profiles are organized in temporal order, with the first observation on the bottom. The vertical dashed line in the left-hand panels indicates the mean radial velocity for HD 156424A. Note that the Stokes V profiles track the RV variation of Stokes I. Note also the clear signatures in several of the N profiles. The vertical dashed lines in the left-hand panels indicate the limits of the Stokes I profile of HD 156424B. Note the lack of RV variation in either Stokes I or V. Asymmetry in Stokes I does not appear to be related in an obvious way to RV variation of HD 156424A’s Stokes I profiles, suggesting that the line profile variability of HD 156424B is not a consequence of contamination by the Stokes I profile of HD 156424A.](https://academic.oup.com/mnras/article-abstract/504/4/4850/6207949?search affliction=global)
Figure 9. Bottom: Hα profiles at maximum (black) and minimum (grey) emission. The solid red and dashed blue show the respective synthetic binary spectra. The vertical lines: dashed black: ±vsini for HD 156424B; dotted blue: Kepler corotation radius RK for HD 156424B; solid purple: maximum emission; dotted green: zero emission. Top: Residual flux after subtraction of synthetic Hα profiles. The top horizontal scale uses vsini for HD 156424B to obtain projected radius from velocity.

about 8× brighter than HD 156424B, the flux modulation due to chemical spots would decrease to about 0.0125–0.125 per cent, or around 0.15–1.5 mmag. The pre-whitened TESS light curve has an upper limit of around 0.15 mmag near 2 c/d, making it possible that the contribution from HD 156424B’s rotational modulation remains hidden in the noise. Note that the star’s relatively low vsini and very rapid rotation indicate that the rotational axis must be nearly aligned with the line of sight, in which case the photometric variation due to chemical spots should be very low in amplitude.

The EWs are shown phased with the 0.52 d period in Fig. 12, using \(T_0 = 2456126.21(4)\) as determined from the maximum emission strength of a harmonic fit. Dynamic spectra of Hα and He i 6678 are shown in Fig. 13. Emission is detectable at all phases, with only weak rotational modulation. This suggests that the obliquity of the magnetic field must be relatively small, with magnetically confined plasma almost equally distributed within the magnetic equatorial plane (e.g. Townsend & Owocki 2005). The low level of variability in Hα also suggests that the light-curve modulation due to rotational modulation of HD 156424B is likely to be fairly modest, consistent with the failure to detect it in the TESS data.

8 MAGNETIC MODELS

8.1 HD 156424A

Since there is no sign of line profile variability coherent with the 1.39 d period identified in TESS photometry, this period probably does not reflect rotational modulation. It is furthermore doubtful that the \(B_r\) measurements can be relied upon to determine the rotational period, as they likely remain contaminated by polarization from HD 156424B. Therefore, the rotational period of this star cannot be determined.
comparable to $\sigma_i$. This indicates that either strongly variable, in all cases presenting a typical Zeeman s-curve.

The purple curve shows $3\times 10^2$ (bottom). Light blue shows periodograms after pre-whitening with the 0.52 d period. The dashed red line shows the 1.4 d period determined from TESS data. The purple curve shows $3\times$ the noise level.

Figure 12. Top to bottom: $\langle B_z \rangle$, H$\alpha$ EW, He i 6678 EW, and Si ii 6347 EW, phased with the rotation period of HD 156424B. EW curves show harmonic fits; shaded regions show fit uncertainties. In the top panel, the solid and dotted lines indicate the mean ($B_z$), and the standard deviation, which is comparable to $\sigma_B$. The red curve shows the inferred magnetic model.

Examination of the Stokes V profiles shows that they are not strongly variable, in all cases presenting a typical Zeeman s-curve. This indicates that either $i$ or $\beta$, or more probably both, are small. First, the low $\sin i$ suggests either small $i$, very slow rotation, or both. There is no indication of long-term evolution over the $\sim 2$ yr of the spectropolarimetric data set, arguing against slow rotation (unless the rotation period is exceptionally long, i.e. several decades). Secondly, a small $i$ and large $\beta$, or large $i$ and small $\beta$, would tend to produce cross-over signatures, which are not detected.

HD 156424A’s magnetic field and magnetospheric parameters were therefore determined under the conservative assumption that $i < 20^\circ$, chosen to be small enough to reflect something close to the true value, yet large enough to permit reasonable coverage of the angular parameter space. The Monte Carlo Hertzsprung–Russell diagram sampler described by Shultz et al. (2019b) was utilized, with the parameter space constrained by the star’s membership in the Sco OB4 association (main-sequence turn-off age $\log t/yr = 6.8 \pm 0.2$, Kharchenko et al. 2005). The resulting parameters are given in Table 3. Since a fit could not be performed to ($B_z$) without a rotation period, parameters for HD156424B were determined in the same fashion as for HD156424A, and are given in Table 3. We proceed on the assumption that the 0.52 d period identified from the vertical offset and semi-amplitude of the sinusoidal fit were approximated with the mean and standard deviation of ($B_z$). This results in $B < 20^2$ and $B_z > 3$ kG, with $2\sigma$ and $3\sigma$ upper limits of 6.2 and 16.2 kG.

The Alfvén radius $R_A$ was calculated in the same fashion as by Shultz et al. (2019b), i.e. using the scaling relationships provided by ud-Doula & Owocki (2002), ud-Doula, Owocki & Townsend (2008), and with the mass-loss rate $M$ and wind terminal velocity $v_{\infty}$ calculated using the Vink, de Koter & Lamers (2001) recipe. The upper limit on the Kepler corotation radius $R_K$ and the lower limit on $R_A$ suggest that the star probably has a small CM ($\log R_A/R_K < 0.4$), but the strength of the magnetic field at the Kepler radius $log B_K/G < 0.8$ is much lower than the threshold for emission determined empirically by Shultz et al. (2020).

While the lack of a rotation period means that only a lower limit can be determined for $R_A > 12 R_*$, given the low $\sin i$ of HD 156424A this lower limit translates to a projected velocity of $60$ km s$^{-1}$, almost $10\times$ less than the maximum extent of H$\alpha$ emission (Fig. 9). HD 156424A would need to have $R_A > 100 R_*$ in order for the H$\alpha$ emission to plausibly belong to this star. Since $R_A$ scales approximately as $B_d^2$ (ud-Doula & Owocki 2002; ud-Doula et al. 2008), this would require $B_d > 200$ kG, almost $10\times$ stronger than the strongest magnetic field seen in an early-type star (Shultz et al. 2019b). Indeed, such a strong magnetic field would produce obvious Zeeman splitting in the star’s spectral lines, which is not observed. It can therefore be firmly excluded that the H$\alpha$ emission originates in a CM around HD 156424A.

### 8.2 HD 156424B

We proceed on the assumption that the 0.52 d period identified from H$\alpha$ is the rotational period of the secondary. The parameter space constrained with the main-sequence turn-off age of Sco OB4. Since a fit to ($B_z$) cannot be obtained, as with HD 156424A we initially approximated the fitting parameters $B_0$ and $B_1$ (see Shultz et al. 2018b) from the mean and standard deviation of ($B_z$), respectively. However, the large standard deviation in ($B_z$), resulting in a geometrical parameter $r = \cos (i + \beta)\cos (i - \beta) = (B_0 - B_1)/(B_0 + B_1) = 0.6 \pm 0.5$ (Preston 1967, 1974), resulted in a maximum likelihood value of $\beta$ that did not give a good reproduction of the actual measurements. The standard deviation in ($B_z$) is less than the mean error bar, and therefore probably overestimates the actual underlying variation in ($B_z$). We instead adopted the mean weighted uncertainty in ($B_z$). 0.1 kG, as both the semi-amplitude $B_1$ and the uncertainty in the semi-amplitude, yielding $r = 0.88 \pm 0.12$. The resulting dipolar ORM model is compared to the ($B_z$) measurements in the top panel of Fig. 12.
the relatively low level of variation in the EWs and although with a large error bar. The very small rotation means that the star must have a non-negligible oblateness (v sin i < 0.1), and a very small Kepler radius (1.6 R_⊙). This in turn means that log B_K/G = 3.0, which is close to the upper extreme of the sample examined by Shultz et al. (2020).

The peak Hα emission strength is about 0.08 nm. In the vicinity of Hα, the synthetic BSTAR2006 spectra used to measure the EWs indicate that HD 156424B should be about 20 per cent as bright as the primary, thus the EW should be scaled by a factor of 5 for an intrinsic emission strength of 0.4 nm. Comparing this to fig. 5 by Shultz et al. (2020), this is almost exactly the predicted emission strength for a star with HD 156424A’s extreme value of B_K. The analytical scaling relationship for Hα emission EW from CMs based on centrifugal breakout, given by Owoki et al. (2020), yields 0.36 ± 0.08 nm, again corresponding very closely to the measured value once corrected for dilution by the light of the primary.

While the emission strength is consistent with B_K, the extent of emission is quite surprising. In all of the stars examined by Shultz et al. (2020), the radius at which emission is at a maximum is between 1 and 2 × R_K. In this case, assuming a linear mapping between radius r and velocity v such that r/R_* = v/v sin i and with v sin i = 25 km s^{-1}, the emission peak occurs at about r_{max} ~ 10 R_*, which is several times larger than R_K. The extent of emission is also extraordinary: about 20 R_*, as compared to the next-most-extensive magnetospheric emission profile, that of CPD–62°2124, about 15 R_*. For r_{max} to occur between 1 and 2 × R_K, v sin i would need to be between about 100 and 200 km s^{-1}. Such high values of v sin i are clearly excluded by the data. If the rotation period is actually twice as long, i.e. 1.04 d, the Kepler radius is about 2.8 R_⊙, and the emission peak then occurs at about 4.4 R_K – still significantly further from R_K than is generally seen.

### Table 3. Rotational, magnetic, and magnetospheric parameters. Values for B_K for HD 156424A correspond to the lower limit and the 3σ upper limit.

| Parameter               | A       | B       |
|-------------------------|---------|---------|
| R_(c) (R_⊙)             | 3.8 ± 0.2 | 2.5 ± 0.1 |
| M_(c) (M_⊙)             | 8.8 ± 0.6 | 4.8 ± 0.4 |
| P_(c) (d)               | 0.52±0.9(9) | – |
| T_(c) (HJD)             | 2456126.22(4) | – |
| v sin i (km s^{-1})     | 5 ± 1 | 25 ± 2 |
| v max (km s^{-1})       | 14 ± 1 | – |
| i (°)                   | <20 | 6 ± 1 |
| v_(Hβ) (kms)            | >6 | 270^{+3.15}_{-0.03} |
| W                       | >0.01 | 0.47^{+0.03}_{-0.01} |
| R_0/R_⊙                | 1 | 0.91^{+0.003}_{-0.001} |
| R_K (R_⊙)               | <10 | 1.65 ± 0.04 |
| B_K (kG)                | 0.82 ± 0.05 | 1.6 ± 0.1 |
| B_1 (kG)                | 0.06 ± 0.05 | 0.4 ± 0.6 |
| β  (°)                  | <20 | 31^{+20}_{-13} |
| B_(4G) (kG)             | >3/4 <16.2 | 8^{+1}_{-5} |
| log (M/M_⊙/yr)          | −9.1^{+0.2}_{-0.4} | −10.7 ± 0.2 |
| v_(nc) (km s^{-1})      | 11000^{+600}_{-50} | 11000^{+70}_{-5} |
| log η_0                 | 4.3 ± 0.3 | 6.8 ± 0.5 |
| R_A (R_⊙)               | >12 | 43^{+2}_{-3} |
| log R_A/R_K              | <0.4 | 1.43 ± 0.05 |
| log (B_K/G)             | <0.8 | 3.0 ± 0.1 |

The result is that the inclination i ~ 6° is very small, as is β ~ 30°, although with a large error bar. The very small i is consistent with the relatively low level of variation in the EWs and \(B_0\). The small \(\beta\) is likewise consistent with the minimal variation in the LSD Stokes V profiles. The surface dipole strength \(B_0 \sim 8 \text{ kG}\). The very rapid rotation means that the star must have a non-negligible oblateness \((R_p/R_c \sim 0.91)\), and a very small Kepler radius \((1.6 R_\odot)\). This in turn means that log \(B_K/G = 3.0\), which is close to the upper extreme of the sample examined by Shultz et al. (2020).

9 DISCUSSION AND CONCLUSIONS

We have analysed the TESS data and the combined high-resolution spectropolarimetric and spectroscopic data set for HD 156424. Several frequencies are detectable in the TESS light curve, four with...
amplitudes of 2–3 mmag, the remainder with amplitudes of order 0.1 mmag; 6 of the 11 detected frequencies are independent, the remainder being harmonics or linear combinations. The majority of the frequencies are above 5 c/d, indicating that HD 156424 is a β Cep pulsator, although the second strongest frequency at about 0.72 d\(^{-1}\) could be indicative of slowly pulsating B-star pulsation. Analysis of RVs also detects the strongest frequency, confirming that these oscillations belong to HD 156424 and not a background star. RVs also show evidence for long-term RV variations with an amplitude of ≈5 km s\(^{-1}\) and a period of order years. The small but statistically significant difference in frequencies obtained from the TESS and RV data sets (separated by several years) is consistent with the light-time effect from orbital motion with the observed RV amplitude.

Close examination of the mean spectrum has revealed the presence of a companion star, HD 156424B. This star is not an RV variable and is therefore not the companion responsible for HD 156424A’s orbital RV variation; instead, we identify it as the orbital companion detected via speckle observations by Hartkopf et al. (1993) and Tokovinin et al. (2010). In an attempt to remove the contribution of HD 156424B’s spectrum from the LSD mean line profile and thereby obtain cleaner magnetic measurements, we found that HD 156424 is itself a magnetic star. This makes HD 156424 the second known doubly magnetic hot binary although, unlike ε Lupi, the stars are not interacting (Shultz et al. 2015b; Pablo et al. 2019). Neither HD 156424A nor HD 156424B have strongly variable Stokes V profiles. The mean (\(B_\gamma\)) for HD 156424A is about −800 G, while that of HD 156424B is about +1.5 kG. In neither case could a rotational period be determined from (\(B_\gamma\)), likely due to a combination of residual contamination of Stokes V by the other star, combined with the overall low level of variation.

The spectrum displays variable H\(\alpha\) emission with a morphology consistent with an origin in a CM; that is, with two emission peaks at velocities greater than \(\sin i\) \(v\) He I, O I, Si II, and Fe II lines are also variable. The variability in Si II is clearly due to HD 156424B. This star is not an RV variable and is therefore not the companion responsible for HD 156424A’s orbital motion with the observed RV amplitude. The variability in Si II is probably not due to rotation. Adjusting for dilution of the spectrum by HD 156424A, the H\(\alpha\) emission is comparable to the strongest yet detected in a magnetic B-type star. This very strong emission is consistent with HD 156424B’s apparently rapid rotation and strong magnetic field. However, peak emission appears at about a distance of 10 \(R_\odot\) from the star, very far above the Kepler radius. This is surprising since in every other case emission strength peaks between 1 and 2 \(R_\odot\). The maximum extent of emission is furthermore about 20 \(R_\odot\), indicating the star’s magnetosphere is very extended compared to similar systems.

**ACKNOWLEDGEMENTS**

This work is based on observations obtained at the CFHT, which is operated by the National Research Council Canada, the Institut National des Sciences de l’Univers of the Centre National de la Recherche Scientifique (INSU,CNRS) of France, and the University of Hawaii; and at the La Silla Observatory, ESO Chile with the 3.6-m telescope and the MPA 2.2-m telescope under programme IDs 187.D-0917(C), 095.D-0269(A), and 095.A-9007. This work has used the VALD data base, operated at Upsala University, the Institute of Astronomy Ras in Moscow, and the University of Vienna. This research has used the SIMBAD data base, operated at CDS, Strasbourg, France. Some of the data presented in this paper were obtained from the MAST. STScI is operated by the Association of Universities for Research in Astronomy, Inc., under NASA contract NAS5-2655. MES acknowledges support from the Annie Jump Cannon Fellowship, supported by the University of Delaware and endowed by the Mount Cuba Astronomical Observatory. ADU acknowledges support from the NSERC Postdoctoral Fellowship Program. GAW acknowledges support from the Natural Sciences and Engineering Research Council (NSERC) of Canada in the form of a Discovery Grant. OK acknowledges support by the Swedish Research Council and the Swedish National Space Agency. The MiMeS collaboration acknowledge financial support from the Programme National de Physique Stellaire of INSU,CNRS. We acknowledge the Canadian Astronomy Data Centre (CADC). The authors thank the anonymous referee for their helpful criticism.

Based on observations obtained at the CFHT, which is operated by the National Research Council of Canada, the INSU of the CNRS of France, and the University of Hawaii; and at the La Silla Observatory, ESO Chile with the MPA 2.2-m telescope.

**DATA AVAILABILITY STATEMENT**

Reduced ESPaDOnS spectra are available at the CADC at https://www.cadc-ccda.hia-iha.nrc-cnrc.gc.ca/en/, and FEROS and HARPSpol spectra are available in raw form at the ESO archive at http://archive.eso.org/eso/esoarchive_main.html. TESS data are available at the MAST archive at https://mast.stsci.edu/portal/Mashup/Clients/Mast/Hosted/MastPortal.html. Data in all archives can be found via standard stellar designations. Reduced ESO data are available from the authors on request.

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APPENDIX: RADIAL VELOCITIES

Table A1. Radial velocity (RV) measurements.

| Instrument    | HD   | RV (km s$^{-1}$) |
|---------------|------|-----------------|
| HARPSpol      | 126.68401 | 1.6 ± 0.4       |
| HARPSpol      | 126.65882 | −0.8 ± 0.7      |
| HARPSpol      | 126.66962 | −2.7 ± 0.4      |
| HARPSpol      | 126.68042 | −2.9 ± 0.3      |
| HARPSpol      | 126.69123 | −1.0 ± 0.4      |
| HARPSpol      | 126.70204 | 1.2 ± 0.7       |
| HARPSpol      | 126.71284 | 3.5 ± 1.1       |
| HARPSpol      | 126.72363 | 3.8 ± 0.7       |
| HARPSpol      | 127.76099 | −1.2 ± 1.1      |
| HARPSpol      | 127.77180 | 0.7 ± 0.3       |
| HARPSpol      | 127.78261 | 3.3 ± 0.3       |
| HARPSpol      | 127.79342 | 4.0 ± 0.8       |
| ESPaDOnS      | 758.03027 | 12.3 ± 0.8      |
| ESPaDOnS      | 758.03596 | 11.5 ± 1.1      |
| ESPaDOnS      | 758.04166 | 10.3 ± 1.0      |
| ESPaDOnS      | 758.04734 | 9.4 ± 1.1       |
| ESPaDOnS      | 761.94523 | 9.1 ± 1.4       |
| ESPaDOnS      | 761.95091 | 10.7 ± 1.6      |
| ESPaDOnS      | 761.95661 | 11.3 ± 0.7      |
| ESPaDOnS      | 761.96230 | 10.2 ± 0.7      |
| ESPaDOnS      | 761.96893 | 9.9 ± 2.0       |
| ESPaDOnS      | 761.97461 | 10.1 ± 0.8      |
| ESPaDOnS      | 761.98029 | 7.0 ± 0.8       |
| ESPaDOnS      | 761.98598 | 5.8 ± 1.0       |
| ESPaDOnS      | 814.95607 | 10.6 ± 1.2      |
| ESPaDOnS      | 814.96176 | 11.4 ± 0.8      |
| ESPaDOnS      | 814.96745 | 10.4 ± 0.8      |
| ESPaDOnS      | 814.97313 | 10.8 ± 0.8      |
| ESPaDOnS      | 814.97908 | 9.0 ± 1.0       |
| ESPaDOnS      | 814.98477 | 7.6 ± 0.9       |
| ESPaDOnS      | 814.99046 | 5.7 ± 1.0       |
| ESPaDOnS      | 814.99615 | 4.3 ± 1.1       |
| ESPaDOnS      | 821.93185 | 10.0 ± 0.9      |
| ESPaDOnS      | 821.93754 | 9.5 ± 1.1       |
| ESPaDOnS      | 821.94323 | 7.8 ± 1.3       |
| ESPaDOnS      | 821.94892 | 6.3 ± 1.1       |
| ESPaDOnS      | 821.97258 | 4.5 ± 1.1       |
| ESPaDOnS      | 821.97827 | 4.8 ± 1.2       |
| ESPaDOnS      | 821.98395 | 5.8 ± 1.0       |
| ESPaDOnS      | 821.98964 | 7.4 ± 0.9       |
| ESPaDOnS      | 824.88691 | 8.7 ± 0.6       |
| ESPaDOnS      | 824.89260 | 7.2 ± 0.7       |
| ESPaDOnS      | 824.89829 | 4.8 ± 1.2       |
Table A1 – continued

| Instrument  | HJD−2456000 | RV (km s$^{-1}$) |
|-------------|-------------|-----------------|
| ESPaDOnS    | 824.90398   | 3.6 ± 1.1       |
| ESPaDOnS    | 824.90983   | 3.0 ± 1.4       |
| ESPaDOnS    | 824.91552   | 3.3 ± 0.6       |
| ESPaDOnS    | 824.92121   | 3.8 ± 0.9       |
| ESPaDOnS    | 824.92690   | 5.2 ± 0.9       |
| FEROS       | 1203.63038  | 8.9 ± 1.0       |
| FEROS       | 1203.69126  | 2.9 ± 0.8       |
| FEROS       | 1204.67172  | 3.5 ± 0.9       |
| FEROS       | 1204.73291  | 7.7 ± 0.9       |
| FEROS       | 1205.49821  | 6.5 ± 1.0       |
| FEROS       | 1205.66785  | 4.5 ± 0.5       |
| FEROS       | 1205.69098  | 9.5 ± 0.6       |
| FEROS       | 1206.50130  | 10.2 ± 1.3      |
| FEROS       | 1206.70782  | 4.7 ± 1.0       |
| FEROS       | 1207.55325  | 8.6 ± 1.3       |
| FEROS       | 1207.71807  | 6.0 ± 0.9       |
| FEROS       | 1208.55215  | 9.2 ± 1.1       |

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