Gravity-driven Lyα blobs from cold streams into galaxies

Tobias Goerdt, A. Dekel, A. Sternberg, D. Ceverino, R. Teyssier and J. R. Primack

1 Racah Institute of Physics, The Hebrew University, Jerusalem 91904, Israel
2 The Raymond and Beverly Sackler School of Physics and Astronomy, Tel Aviv University, Tel Aviv 69978, Israel
3 Institut für Theoretische Physik, Universität Zürich, Winterthurer Strasse 190, CH-8057 Zürich, Switzerland
4 Department of Physics, University of California, Santa Cruz, CA 95064, USA

Accepted 2010 April 28. Received 2010 April 26; in original form 2009 November 30

ABSTRACT
We use high-resolution cosmological hydrodynamical adaptive mesh refinement (AMR) simulations to predict the characteristics of Lyα emission from the cold gas streams that fed galaxies in massive haloes at high redshift. The Lyα luminosity in our simulations is powered by the release of gravitational energy as gas flows from the intergalactic medium into the halo potential wells. The ultraviolet UV background contributes only $< 20\%$ to the gas heating. The Lyα emissivity is due primarily to electron-impact excitation cooling radiation in gas at $\sim 2 \times 10^4$ K. We calculate the Lyα emissivities assuming collisional ionization equilibrium at all gas temperatures. The simulated streams are self-shielded against the UV background, so photoionization and recombination contribute negligibly to the Lyα line formation. We produce theoretical maps of the Lyα surface brightnesses, assuming that $\sim 85\%$ of the Lyα photons are directly observable. We do not consider transfer of the Lyα radiation, nor do we include the possible effects of internal sources of photoionization such as star-forming regions. Dust absorption is expected to obscure a small fraction of the luminosity in the streams. We find that typical haloes of mass $M_v \sim 10^{12} - 10^{13} M_\odot$ at $z \sim 3$ emit as Lyα blobs (LABs) with luminosities $10^{43} - 10^{44}$ erg s$^{-1}$. Most of the Lyα comes from the extended ($50 - 100$ kpc) narrow, partly clumpy, inflowing, cold streams of $(1 - 5) \times 10^4$ K that feed the growing galaxies. The predicted LAB morphology is therefore irregular, with dense clumps and elongated extensions. The integrated area contained within surface brightness isophotes of $2 \times 10^{-18}$ erg s$^{-1}$ cm$^{-2}$ arcsec$^{-2}$ is $\sim 2 - 100$ arcsec$^2$, consistent with observations. The linewidth is expected to range from $10^2$ to more than $10^3$ km s$^{-1}$ with a large variance. The typical Lyα surface brightness profile is $\propto r^{-1.2}$, where $r$ is the distance from the halo centre. Our simulated LABs are similar in luminosity, morphology and extent to the observed LABs, with distinct kinematic features. The predicted Lyα luminosity function is consistent with observations, and the predicted areas and linewidths roughly recover the observed scaling relations. This mechanism for producing LABs appears inevitable in many high-$z$ galaxies, though it may work in parallel with other mechanisms. Some of the LABs may thus be regarded as direct detections of the cold streams that drove galaxy evolution at high $z$.

Key words: galaxies: evolution – galaxies: formation – galaxies: high-redshift – intergalactic medium – galaxies: ISM – cosmology: theory.

1 INTRODUCTION
Hundreds of Lyα blobs (LABs) have been detected so far in the redshift range $z = 2 - 6.5$, mostly near $z \sim 3$ (Steidel et al. 2000; Matsuda et al. 2004; Saito et al. 2006; Yang et al. 2008; Ouchi et al. 2009). Their luminosities range from below $10^{43}$ to above $10^{44}$ erg s$^{-1}$, and they extend on the sky to $30 - 50$ kpc.
and more. The two main open questions are (i) the origin of the extended cold and relatively dense gas capable of emitting Lyα and (ii) the continuous energy source for exciting the gas to emit Lyα.

The emitting hydrogen gas should be at a temperature $T \gtrsim 10^4$ K, relatively dense and span a much larger area than covered by the stellar component of galaxies. It may arise from outflows or inflows. The energy sources discussed in the literature include photoionization by obscured active galactic nuclei (AGNs), early starbursts or extended X-ray emission (Haiman & Rees 2001; Scharf et al. 2003; Jimenez & Haiman 2006), as well as compression of ambient gas by superwinds to dense Lyα emitting shells (Mori, Umemura & Ferrara 2004), and star formation triggered by relativistic jets from AGNs (Rees 1989).

Many of the bright LABs are found in the vicinity of massive, star-forming galaxies (Matsuda et al. 2006). Multiwavelength observations reveal that a fraction of the LABs are associated with submillimetre and infrared sources that indicate very high star formation rates (SFRs) in the range $10^8 - 10^9 \, M_\odot \, yr^{-1}$ (Chapman et al. 2001; Geach et al. 2005, 2007) or with obscured AGNs (Basu-Zych & Scharf 2004; Scarlata et al. 2009). While stellar feedback and AGNs could in principle provide the energy source for the Lyα luminosity, many LABs are not associated with sources of this sort that are powerful enough to explain the observed Lyα luminosities (Smith & Jarvis 1999; Nilson et al. 2006). This indicates that star formation and AGNs are not the sole drivers of LABs, and may not even be the dominant source for LABs.

Indeed, high-redshift galaxies exhibit a generic mechanism that simultaneously provides both the cold gas and the energy source. It is a direct result of the phenomenon robustly established by simulations and theoretical analysis, where high-$z$ massive galaxies are continuously fed by narrow, cold, intense, partly clumpy, gaseous streams that penetrate through the shock-heated halo gas into the inner galaxy, grow a dense, unstable, turbulent disc with a bulge and trigger rapid star formation (Birnboim & Dekel 2003; Keres et al. 2005; Dekel & Birnboim 2006; Ocvirk, Pichon & Teyssier 2008; Dekel et al. 2009a; Dekel, Sari & Ceverino 2009b; Johansson, Naab & Ostriker 2009; Keres et al. 2009; Ceverino, Dekel & Bournaud 2010). Massive clumpy star-forming discs observed at $z \approx 2$ (Genel et al. 2008; Genzel et al. 2008; Förster Schreiber et al. 2009) may have been formed via the smooth and steady accretion provided by cold flows, as opposed to merger events. The streaming of the gas into the dark matter halo potential well is associated with transfer of gravitational energy to excitations of the hydrogen atoms followed by cooling emission of Lyα (Haiman, Spaans & Quataert 2000; Fardal et al. 2001; Furlanetto et al. 2003, 2005; Dekel & Birnboim 2006, 2008; Khochfar & Ostriker 2008).

There were two earlier attempts to compute the Lyα cooling radiation from hydrodynamically smoothed particle hydrodynamics (SPH) cosmological simulations. Based on their analysis, Fardal et al. (2001) pointed out the potential association of this Lyα cooling emission with the first observed LABs. These early simulations did not allow a proper resolution of the detailed structure of the cold streams as they penetrate through the hot medium. Their shortcomings included intrinsic limitations of the SPH technique, a limited force resolution of $7 \, h^{-1} \, kpc$ (comoving), not allowing for radiative cooling below $10^4$ K and neglecting the photoionization by the ultraviolet (UV) background. Furlanetto et al. (2005) used SPH simulations with a somewhat higher resolution of $\sim 1 \, kpc$ to make more detailed comparisons of the simulated and observed luminosity functions and size distributions of LABs powered by cold accretion. In their pessimistic model, assuming no emissivity from gas that is self-shielded from the ionizing background radiation, they end up with low Lyα luminosity. They comment that cooling intergalactic medium (IGM) gas may explain the observations only if one adopts an optimistic scenario, where the self-shielded gas is emitting Lyα at collisional ionization equilibrium (CIE). The latter scenario will also be adopted in this paper. Other sources of ionization are ignored and radiative transfer is not applied.

Dijkstra & Loeb (2009, hereafter DL09) have recently worked out an analytic toy model for Lyα cooling radiation from cold streams, based on the general properties of the cold streams as reported from cosmological simulations. They conclude that the streams could in principle provide spatially extended Lyα sources with luminosities, linewidths and abundances that are similar to those of observed LABs. They point out that the filamentary structure of cold flows may explain the wide range of observed LAB morphologies. They also highlight the fact that the most luminous cold flows are associated with massive haloes, which preferentially reside in dense large-scale surroundings, in agreement with the observed presence of bright LABs in dense environments. This model presents a successful feasibility test for the role of cold streams in powering the LAB emission, and it provides physical intuition into the way by which this process could be manifested. However, a comparison to our simulations indicates that this simplified model does not capture the detailed hydrodynamic properties of the cold streams. In particular, it significantly overpredicts the cold gas density, underpredicts the gas temperature, and does not account for the partly clumpy nature of the streams and their characteristic radial distribution in the halo. As a result, the DL09 model underestimates the total Lyα luminosity by a factor of a few, and it therefore has to appeal to the excessive clustering of the LABs in order to boost up the predicted luminosity function for a match with the observations.

In this paper, we calculate the Lyα emission from the cold gas in high-redshift galactic haloes using state-of-the-art hydrodynamical adaptive refinement tree (AMR) cosmological simulations. With a maximum resolution better than 70 pc in our simulations we can quite accurately map the extended Lyα sources. This allows us to study their individual shapes, morphologies and kinematics. We measure quantities such as the distribution of surface brightness within each halo, the area covered with surface brightness above a given isophotal threshold, the predicted observed linewidth, the typical total Lyα luminosity per halo of a given mass, and the overall Lyα luminosity function of LABs. These predicted properties are compared with the observed LABs.

This paper is organized as follows. In Section 2 we work out a simple feasibility test where we use a simple toy model to estimate the expected gravitational heating power as a proxy for the Lyα luminosity. In Section 3 we introduce the two sets of cosmological simulations used. In Section 4 we explain our methodology of computing Lyα emissivity and luminosity as a function of halo mass. In Section 5 we identify the gas that contributes to the Lyα emission. In Section 6 we apply our methodology to the simulations and provide predicted images of LABs. In Section 7 we determine the scaling of Lyα luminosity as a function of halo mass and redshift. In Section 8 we compare our predicted Lyα luminosity function with observational results. In Section 9 we measure the predicted surface density profile, isophotal area and linewidth, and compare to observations. In Section 10 we show that gravitational heating is the main source of energy driving the Lyα luminosity in our simulations, while the role of photoionization is minor. In Section 11 we discuss our analysis and results and draw our conclusions.
2 FEASIBILITY OF GRAVITATIONAL HEATING

The gravitational energy gain due to the streaming of the gas from the virial radius towards the centre of the halo potential well is a natural source of energy for the Ly$\alpha$ emission. The gravitational energy released per unit infalling mass is of the order of $V^2 = GM_c/r_c$, where the quantities are the virial mass, and radius, respectively. The accretion rate $M$ can be estimated from the observed SFR. For $V \sim 300\,\text{km\,s}^{-1}$ and $M \sim 150\,\text{M}_\odot\,\text{yr}^{-1}$ we obtain a power $\dot{M} V^2 \sim 10^{41}\,\text{erg\,s}^{-1}$, comparable to the luminosity of a typical LAB.

In more detail, an analysis of the MareNostrum (MN) cosmological simulation (described in Section 3) reveals that both the cold gas accretion rate $\dot{M}_c$ and its inward velocity are roughly constant along the streams (Dekel et al. 2009a). If so, the gravitational power deposited at radius $r$ per unit radial length in the cold gas is

$$\dot{E}_{\text{grav}} = f_c \dot{M}_c \left| \frac{\partial \phi}{\partial r} \right|, \quad (1)$$

where $\phi(r)$ is the gravitational potential at $r$. The factor $f_c$ is the fraction of the energy that goes to heating the cold streams themselves rather than the hot medium. The total power deposited between the virial radius $R_v$ and radius $r$, across a potential difference $\Delta \phi(R_v, r)$, is

$$\dot{E}_{\text{grav}} = f_c \dot{M}_c \left| \Delta \phi(R_v, r) \right|. \quad (2)$$

Given the total mass density profile $\rho(r)$ and the associated mass profile $M(r)$, the potential gain is

$$|\Delta \phi| \equiv \phi V^2, \quad \phi = -1 + \frac{V^2(r)}{V_c^2} + \int_{r/R_v}^{1} \frac{3\rho(r')}{\rho_v} \, dr', \quad (3)$$

where $V^2(r) = GM(r)/r$, $\rho_v$ is the mean density within the virial radius and $\phi$ is typically a number of the order of unity.

For an NFW potential with a concentration parameter $C$ (Navarro, Frenk & White 1997), one has

$$\phi(r) = \frac{C}{A_1(C)} \left\{ \ln(1 + x) - \frac{\ln(1 + C)}{C} \right\}, \quad (4)$$

where $x = Cr/R_v$ and $A_1(x) = \ln(x + 1) - x/(x + 1)$. According to the virial condition and sharpness of the N-body simulations (Bullock et al. 2001), the average concentration parameter can be a function of halo mass and redshift is $C \sim 3M_{12}^{-0.13} \left(1 + z\right)^{-1}$, where $M_{12} = M_c/10^{12}\,\text{M}_\odot$ and $(1 + z) = 1 + 0.25$. For $M_c \sim 10^{12}\,\text{M}_\odot$, this is $C \approx 3$ at $z = 3$. Then $\phi(r) \approx 4.7\ln(1 + x)/x - 0.46 \approx 4.7\ln(1 + x)/x \approx 2.5$. A practical approximation for the average accretion rate into haloes of $M_c \sim 10^{12} - 10^{13}\,\text{M}_\odot$ in the standard Lambda cold dark matter (CDM) cosmology is derived from the Extended Press–Schechter approximation and from the Millennium cosmological simulation (Neistein, van den Bosch & Dekel 2006; Binomak, Dekel & Neistein 2007; Neistein & Dekel 2008, appendix A). The number density at redshifts 2–3 (Förster Schreiber et al. 2006; Genzel et al. 2006; Elmegreen et al. 2007; Genzel et al. 2008; Stark et al. 2008).

The virial velocity is given by

$$V_c \simeq 236\,\text{km\,s}^{-1} M_{12}^{1/2} \left(1 + z\right)^{1/2}. \quad (6)$$

So we finally obtain from equation (2),

$$E_{\text{grav}} \simeq 1.2 \times 10^{43}\,\text{erg\,s}^{-1} f_c M_{12}^{1/2} \left(1 + z\right)^{2.25}. \quad (7)$$

Assuming that a substantial fraction $f_c$ of this energy is emitted as observable Ly$\alpha$ radiation, we conclude that at $z \approx 3$, with $f_{\alpha}/f_c \approx 1$, a LAB of luminosity $L \sim 10^{43}\,\text{erg\,s}^{-1}$ is feasible from haloes of $M_c \sim 10^{12}\,\text{M}_\odot$. Luminosities as high as $\sim 10^{44}\,\text{erg\,s}^{-1}$ require haloes of $\sim 3 \times 10^{12}\,\text{M}_\odot$. If $f_{\alpha}/f_c$ is only $\lesssim 0.1$, then the required haloes for the same luminosities are about three times more massive. We note that the mean comoving number density for haloes more massive than $(1, 3, 10) \times 10^{12}\,\text{M}_\odot$ at $z = 3$ is $(4, 0.68, 0.057) \times 10^{-4}\,\text{Mpc}^{-3}$, respectively. Given that LABs of a luminosity higher than $10^{43}\,\text{erg\,s}^{-1}$ appear with a comoving number density of $\sim 5 \times 10^{-5}\,\text{Mpc}^{-3}$ (Fig. 14), the simple gravitational heating model indicates that they can emerge from haloes of $\sim 3 \times 10^{12}\,\text{M}_\odot$, with $f_{\alpha}/f_c \sim 0.1$. Thus, the comparison of the estimates from our toy model with the total luminosities of the observed LABs indicates that gravitational heating is a feasible source for the Ly$\alpha$ emission. This kind of energy has to be released from these galaxies.

One can combine equations (4) and (7) to evaluate the gravitational energy deposited at different radii,

$$E_{\text{grav}}(<r) \simeq 1.2 \times 10^{43}\,\text{erg\,s}^{-1} f_c M_{12}^{1/2} \left(1 + z\right)^{2.25} \times \left[ 1.86 - \frac{1.86 \ln(1 + Cr/R_v)}{Cr/R_v} \right]. \quad (8)$$

This gives a very crude estimate for the 3D luminosity profile $L(<r) = f_c f_{\alpha} \dot{E}_{\text{grav}}(<r)$. This profile is shown in Fig. 1, normalized to the total luminosity inside the virial radius. One can read from this plot the fraction of the luminosity that is expected in the different zones of the halo. For example, about half the luminosity is expected from outside $r = 0.27R_v$, implying that in haloes of $M_c \sim 10^{12}\,\text{M}_\odot$ the LABs are expected to extend over more than 50 kpc. The luminosity profile predicted by gravitational heating is to be compared with our results from the simulations (Section 10), and with observed LABs (Section 9).

3 SIMULATIONS

We use here simulated galaxies from two different suites of simulations, both employing Eulerian AMR hydrodynamics in a cosmological setting. One suite consists of three simulations zooming in on the horizon MN simulation containing hundreds of massive galaxies in a cosmological box of side $50\,\text{h}^{-1}\,\text{Mpc}$ with a maximum resolution of $\sim 1\,\text{kpc}$ (Ocvirk et al. 2008, hereafter MN).

Fig. 2 shows sample gas density maps of galaxies from the two suites of simulations. They demonstrate the dominance of typically three, narrow cold streams, which come from well outside the virial radius along the dark matter filaments of the cosmic web, and penetrate into the discs at the halo centres. The streams are partly clumpy.
Figure 1. Cumulative luminosity profile, showing the fraction of the luminosity from within a sphere of radius $r$ compared to the virial sphere. The toy-model prediction (solid red) is proportional to the fraction of gravitational energy that is deposited inside radius $r$ based on equation (8). The predicted profile in the range $(0.2–0.5)R_v$ can be approximated by the power law $L(<r) \propto r^{0.7}$. About half the luminosity is expected to originate from the cold streams in the halo outside the inner 0.27$R_v$ sphere. Shown in comparison is the luminosity profile as derived from the three simulated CDB galaxies stacked together (see Section 10). The similarity between the toy-model predictions and the simulation results is remarkable.

and partly smooth, even in the simulation of higher resolution. The typical densities in the streams are in the range $n = 0.01–0.1$ cm$^{-3}$, and they reach $n = 0.1–1$ cm$^{-3}$ at the clump centres and in the central disc.

3.1 High-resolution ART CDB simulations

The CDB simulations have been run with the AMR tree code ART (Kravtsov, Klypin & Khokhlov 1997; Kravtsov 2003) with a spatial resolution better than 70 pc in physical units. The code incorporates the relevant physical processes for galaxy formation, including gas cooling and photoionization heating, star formation, metal enrichment and stellar feedback (Ceverino & Klypin 2009). Cooling rates were computed for the given gas density, temperature, metallicity and UV background based on CLOUDY (Ferland et al. 1998), assuming cooling at the centre of a cloud of thickness 1 kpc (Ceverino-Rodriguez 2008; Robertson & Kravtsov 2008). Metals are included, assuming an abundance of 0.02 relative to the total mass for a solar composition. The code implements a ‘constant’ feedback model, in which the combined energy from stellar winds and supernova explosions is released as a constant heating rate over 40 Myr, the typical age of the lightest star that explodes as a Type II supernova (SNII). Photoheating is also taken into account self-consistently with radiative cooling. A uniform UV background based on the Haardt & Madau (1996) model is assumed, and local sources are ignored. In order to mimic the self-shielding of dense, galactic neutral hydrogen from the cosmological UV background, the simulation assumes a substantially suppressed UV background ($5.9 \times 10^{26}$ erg s$^{-1}$ cm$^{-2}$ Hz$^{-1}$, the value of the pre-reionization UV background at $z = 8$) for the gas at total gas densities above $n = 0.1$ cm$^{-3}$.

The unique feature of this code for the purpose of simulating the detailed structure of the streams and the gravitational instability in the disc is to allow the gas to cool down well below $10^4$ K, thus reaching high densities in pressure equilibrium with the hotter and more dilute medium. A non-thermal pressure floor has been implemented to ensure that the Jeans length is resolved by at least seven resolution elements and thus prevent artificial fragmentation on the smallest grid scale (Truelove et al. 1997; Robertson & Kravtsov 2008; CDB). It is effective in the dense ($n > 10$ cm$^{-3}$) and cold ($T < 10^4$ K) regions inside galactic discs, while most of the Lyα emitting gas, which is at temperatures $(1–5) \times 10^4$ K, is not affected by this pressure floor.

Figure 2. Gas density in simulated galaxies from CDB (left-hand panel) and MN (right-hand panel). The colour refers to the maximum density along the line of sight. The contours mark $n = 0.1, 0.01$ and 0.001 cm$^{-3}$, respectively. The circle shows the virial radius. Left-hand panel: one of the three CDB galaxies (resolution 70 pc) at $z = 2.3$, with $M_v = 3.5 \times 10^{11}$ M$_\odot$. Right-hand panel: one of the MN galaxies (resolution 1 kpc) at $z = 2.5$, with $M_v = 10^{12}$ M$_\odot$. In both cases, the inflow is dominated by three cold narrow streams that are partly clumpy. The density in the streams is $n = 0.003–0.1$ cm$^{-3}$, with the clump cores reaching $n \sim 1$ cm$^{-3}$.

© 2010 The Authors. Journal compilation © 2010 RAS, MNRAS 407, 613–631

Downloaded from https://academic.oup.com/mnras/article-abstract/407/1/613/986273
by guest on 28 July 2018
The equation of state remains unchanged at all densities. Stars form in cells where the gas temperature is below \( 10^4 \text{K} \) and its density is above the threshold \( n = 1 \text{ cm}\(^{-3}\) \) according to a stochastic model that is roughly consistent with the Kennicutt (1998) law. The interstellar medium (ISM) is enriched by metals from SNeII and SNeIa. Metals are assumed to be released from each stellar particle by SNIa at a constant rate for 40 Myr since its birth, assuming a Miller & Scalo (1979) initial mass function (IMF) and matching the results of Woosley & Weaver (1995). The metal ejection by SNIa assumes an exponentially declining SNIa rate from a maximum at 1 Gyr. The code treats the advection of metals self-consistently (Ceverino-Rodriguez 2008).

The dark matter particle mass is \( 5.5 \times 10^4 \text{M}_\odot \), the minimum star particle mass is \( 10^5 \text{M}_\odot \), the smallest cell size is 70 pc (physical units), and the force softening length is 105 pc.

The initial conditions for the CDB simulations were created using a low-resolution cosmological N-body simulation in a comoving box of side \( 20 h^{-1} \text{Mpc} \), for which the cosmological parameters were motivated by the WMAP5 following values (Komatsu et al. 2009): \( \Omega_m = 0.27, \Omega_\Lambda = 0.73, \Omega_b = 0.045, h = 0.7 \) and \( \sigma_8 = 0.82. \) At \( z = 1 \), three haloes of \( M \lesssim 10^{12} \text{M}_\odot \) each have been selected, avoiding haloes that were subject to a major merger near that time. The three halo masses at \( z = 2.3 \) are 3.5, 4.6 \( \times 10^{11} \text{M}_\odot \), and they end up as \( 3 \times 10^{12} \text{M}_\odot \) haloes today. For each halo, a concentric sphere of radius twice the virial radius was identified for resimulation with high resolution. Gas was added to the box following the dark matter distribution with a fraction \( f_g = 0.15 \). The whole box was then resimulated, with refined resolution only in the selected volume about the respective galaxy.

3.2 RAMSES MareNostrum simulations

The MN simulation uses the AMR code RAMSES (Teyssier 2002). The spatial resolution is \( \sim 1 \text{kpc} \) in physical units. UV heating is included assuming the Haardt & Madau (1996) background model, as in the CDB simulation. The code incorporates a simple model of supernova feedback and metal enrichment using the implementation described in Dubois & Teyssier (2008). The cooling rates are calculated assuming ionization equilibrium for H and He, including both collisional ionization and photoionization (Katz, Hernquist & Weinberg 1992). Metal cooling is also included using tabulated CLOUDY rates, and is assumed proportional to the metallicity, relative to the Grevesse & Sauval (1998) solar abundances. Unlike in the CDB simulation, no cooling below \( T < 10^4 \text{K} \) is computed, and no self-shielding of the UV flux is assumed. Because we assume that the Ly\( \alpha \) emissions are produced in predominantly shielded gas that is in CIE, our MN computations should be regarded as less accurate than the CDB simulation.

For high-density regions, the MN code considers a polytropic equation of state with \( y_0 = 5/3 \) to model the complex, multiphase and turbulent structure of the ISM (Yepes et al. 1997; Springel & Hernquist 1999) in a simplified form (see Schaye & Dalla Vecchia 1999; Dubois & Teyssier 2008). The ISM is defined as gas with hydrogen density greater than \( n_H = 0.1 \text{ cm}\(^{-3}\) \), one order of magnitude lower than in the CDB simulation. Star formation has been included, for ISM gas only, by spawning star particles at a rate consistent with the Kennicutt (1998) law derived from local observations of star-forming galaxies.

The MN simulation implemented a pressure floor in order to prevent artificial fragmentation, by keeping the Jeans length-scale, \( L_J \propto T_n^{2/3} \), larger than four grid cell sides everywhere. In any case where \( n > 0.1 \text{ cm}\(^{-3}\) \), a density-dependent temperature floor was imposed. It mimics the average thermal and turbulent pressure of the multiphase ISM, in the spirit of Springel & Hernquist (1999) or Dalla Vecchia & Schaye (2008). In our case, we allow the gas to heat up above this temperature floor and cool back. The temperature floor follows a polytropic equation of state with \( T_{\text{floor}} = T_0 (n/n_0)^{\gamma-1} \), where \( T_0 = 10^4 \text{K} \) and \( n_0 = 0.1 \text{ atoms cm}^{-3} \). The resulting pressure floor is given by \( P_{\text{floor}} = n_H k_B T_{\text{floor}}. \)

We crudely correct for this artificial temperature boost in post-processing by subtracting \( T_{\text{floor}} \) from the temperature read from the grid cells where \( n > 0.1 \text{ cm}^{-3} \). If the corrected temperature is below \( 10^4 \text{K} \), we set it to \( 10^4 \text{K} \). In practice, almost all cells where \( n > 0.1 \text{ cm}^{-3} \) are set to \( T = 10^4 \text{K} \). As will become clear in Section 4, this means neglecting any Ly\( \alpha \) emission from these high-density cells. We will evaluate the possible error made by this procedure in the MN galaxies using the more accurate high-resolution CDB galaxies, where no temperature floor has been applied.

For each stellar population, 10 per cent of the mass is assumed to turn into SNeII after 10 Myr, where the energy and metals are released in an impulse. For each supernova, 10 per cent of the ejected mass is assumed to be pure metals, with the remaining 90 per cent keeping the metallicity of the star at birth. SNIa feedback has not been considered.

The dark matter particle mass is \( 1.16 \times 10^5 \text{M}_\odot \), the star particle mass is \( 2.05 \times 10^6 \text{M}_\odot \), the smallest cell size is 1.09 kpc physical and the force softening length is 1.65 kpc.

The initial conditions of the MN simulation were constructed assuming a \( \Lambda \)CDM universe with \( \Omega_m = 0.3, \Omega_\Lambda = 0.7, \Omega_b = 0.045, h = 0.7 \) and \( \sigma_8 = 0.9 \) in a periodic box of side \( 50 h^{-1} \text{Mpc} \). The adaptive-resolution rules in this simulation were the same everywhere, with no zoom-in resimulation of individual galaxies.

4 COMPUTING THE Ly\( \alpha \) LUMINOSITY

Given the gas temperature \( T \) and mass density \( \rho \) in every cubic grid cell of the hydrodynamic simulation, we compute the local Ly\( \alpha \) emissivity produced by electron impact excitation of neutral hydrogen,

\[
\epsilon = n_e n_H \, q_{1s \rightarrow 2p}(T) \, h\nu_{\text{Ly}\alpha} \, \text{erg cm}^{-3} \text{s}^{-1}.
\]

In this expression, \( n_e \) and \( n_H \) are the local electron and atomic hydrogen particle densities (in cm\(^{-3}\)) and \( h\nu_{\text{Ly}\alpha} = 10.2 \text{eV} = 1.63 \times 10^{-11} \text{erg} \) is the Ly\( \alpha \) photon energy. The temperature-dependent quantity \( q_{1s \rightarrow 2p} \) is the collisional excitation rate coefficient (Callaway, Unnikrishnan & Oza 1987),

\[
q_{1s \rightarrow 2p} = \frac{2.41 \times 10^{-6}}{T^{0.5}} \left( \frac{T}{10^4} \right)^{0.22} \times \exp \left( -\frac{h\nu_{\text{Ly}\alpha}}{kT} \right) \, \text{cm}^3 \text{s}^{-1},
\]

where \( T \) is in Kelvin. Radiative reemissions of electrons with protons contribute negligibly to the Ly\( \alpha \) emissivity, as can be seen from the green, long-dashed curve in Fig. 3.

We assume a primordial helium mass fraction \( Y = 0.24 \), corresponding to a helium particle abundance of 1/12 relative to hydrogen. This gives

\[
\rho = (4/3) \, m_p n_H.
\]

where \( m_p \) is the hydrogen mass and \( n_H \) is the total density of hydrogen nuclei (neutral plus ionized). The electron and atomic hydrogen densities in equation (9) are given by

\[
n_{\text{H}1}/n_H = x_{\text{H}1}
\]

© 2010 The Authors. Journal compilation © 2010 RAS, MNRAS 407, 613–631
\( n_e/n_H = x_{HI} + (1/12)(x_{HeII} + 2x_{HeIII}) \) \hspace{2cm} (13)

where we adopt the temperature-dependent ionization fractions, \( x_{HI} \), \( x_{HeII} \), and \( x_{HeIII} \) as computed for CIE assuming case-B hydrogen recombination (O. Gnat, private communication; see also Gnat & Sternberg 2007). A large fraction (\( \sim 95 \% \)) of the Ly\( \alpha \) emission is produced in gas where the helium is fully neutral. In this limit, \( x_{HI} = \alpha_R/\alpha_1 \), which are the recombination and collisional ionization rate coefficients (Gnat & Sternberg 2007).

We note for completeness that approximate fits can be provided by

\[
\alpha_B(T) = 4.9 \times 10^{-6} T^{-1.5} \left( 1 + \frac{115}{T_{0.41}} \right)^{-2.24} \text{ cm}^3 \text{ s}^{-1} \]

\hspace{2cm} (15)

and

\[
q_i = 21.11 \text{ cm}^3 \text{ s}^{-1} K^{3/2} T^{-3/2} \exp \left( -\frac{T_{HI}}{T} \right) \times \left( 2T_{HI}/T \right)^{-1.089} \left[ 1 + (5.65 T_{HI}/T)^{0.874}\right]^{-1.01} ,
\]

\hspace{2cm} (16)

with \( T_{HI} = 1.58 \times 10^5 \text{ K} \) being the ionization threshold according to Hui & Gnedin (1997). We do not use these fits in our calculation here.

We assume that the cold streams are practically self-shielded against the photoionizing UV background, for the following reason. The typical H\( \text{I} \) column densities along the shortest dimensions of the streams are \( \sim 10^{20} \text{ cm}^{-2} \) and the total hydrogen particle densities are in the range 0.01–0.1 cm\(^{-3}\), or typically \( \sim 0.03 \text{ cm}^{-3} \) (see Fig. 6). At redshift \( z = 3 \), the mean Lyman continuum photon intensity is \( \sim 4 \pi I^\alpha \approx 2.2 \times 10^5 \text{ photons s}^{-1} \text{ cm}^{-2} \) (Haardt & Madau 1996) and the unattenuated hydrogen photoionization rate is \( \Gamma = 5.6 \times 10^{-13} \text{ s}^{-1} \). Thus, for gas densities \( n \lesssim 2 \Gamma/\alpha_\odot \approx 4.3 \text{ cm}^{-3} \), the gas will be more than 50 per cent ionized by the unattenuated field. However, absorption of the radiation field in the stream gas will produce a photoionized column

\[
N_{\text{HII}}^\text{photo} = \frac{2\pi f J^\star}{n_{\text{HII}} \alpha_\odot} = 4.2 \times 10^{17} \text{ cm}^{-2}.
\]

(17)

For the typical stream volume densities this is small compared to the neutral columns of \( \sim 10^{20} \text{ cm}^{-2} \) that we find, so the streams can be practically assumed to be self-shielded against the background radiation. Parts of the streams may not be self-shielded against local sources of radiation within the streams, e.g. from star-forming satellites. In those cases we might overpredict the H\( \text{I} \) fraction and thus the Ly\( \alpha \) luminosity.

The CIE neutral hydrogen fraction as a function of temperature and the temperature dependence of the Ly\( \alpha \) emissivity for a given total density are shown in Fig. 3. One can see that the gas is neutral at \( T = 10^4 \text{ K} \) and below, and it becomes highly ionized at \( T > 2 \times 10^4 \text{ K} \). The maximum emissivity at a given density is obtained at \( T \approx 1.8 \times 10^4 \text{ K} \), and the emissivity is high enough to substantially contribute to the overall luminosity in the range \( T = (1-5) \times 10^4 \text{ K} \). Hydrogen in this temperature range is thus the expected source of Ly\( \alpha \) emission.

The total Ly\( \alpha \) luminosity from a given volume is assumed to be a constant fraction of the sum over all the cells that are contained in that volume,

\[
L_{\text{Ly} \alpha} = f_\alpha \sum_i \epsilon_i V_i ,
\]

(18)

where \( V_i \) is the cell volume. The complex radiation transfer process along the cosmological line of sight is summed up in the transmission factor \( f_\alpha \), the fraction of the Ly\( \alpha \) photons that make it to the observer. In Section 8, we will determine the actual value of \( f_\alpha \) by matching the predictions from the simulations with a preliminary determination of the observed Ly\( \alpha \) luminosity function. We expect \( f_\alpha \) to be in the range 0.5–1, because the main source of opacity is likely to be intervening as well as lower redshift intergalactic H\( \text{I} \), which is expected to absorb part of the blue side of the line profile. An estimate for a typical line of sight to \( z \sim 3 \) is \( f_\alpha \sim 0.85 \) (e.g. Faucher-Giguere et al. 2008). The transmission factor along the line of sight to a typical LAB could be slightly smaller because the LABs tend to reside in overdense environments. A smooth component of H\( \text{I} \) in the emitting halo may absorb some of the red wing as well. On the other hand, as estimated next, dust opacity is expected to reduce \( f_\alpha \) by a small factor only except in the galaxy itself and in its immediate vicinity.

Given the gas metallicity in every cell of the simulation, we can estimate the dust absorption as follows. For a medium in which the dust abundance scales linearly with the metallicity \( Z \) (with \( Z \approx 1 \) corresponding to Galactic dust), the continuum opacity due to dust absorption at the Ly\( \alpha \) wavelength is \( \tau_{\text{dust}} \approx 0.1 Z N_{20} \), where \( N_{20} \) is the neutral hydrogen column density in units of \( 10^{20} \text{ cm}^{-2} \) (Draine 2003). The line optical depth is \( \tau_{\text{Ly} \alpha} \approx 0.1 N_{20} T_4^{-1/2} \) (Neufeld 1990) so that \( \tau_{\text{dust}} \) is comparable to \( \tau_{\text{Ly} \alpha} \). For Ly\( \alpha \) scattering through a plane parallel medium the total line optical depth traversed is

\[
\tau_{\text{dust}} \approx 2.9 N_{20}^{4/3} T_4^{-1/3}.
\]

(19)
For $Z = 0.1$, assuming $T_A \sim 1$, we obtain $\tau_{\text{dust}} = 1$ for $N_{20} \simeq 3$. We find in the CDB galaxies that the Lyα luminosity-weighted fraction of volume elements that have a column density $N_{20} \geq 3$ between them and the observer, averaged over different directions, is less than a third. This provides an estimate of no more than a third reduction outside $r \sim 0.1 R_e$ (where $Z \lesssim 0.1$).

Fig. 4 shows the stacked 3D metallicity profile of the Lyα emitting gas in the three simulated CDB galaxies. It is computed for gas in the temperature range $(1-5) \times 10^4 \text{ K}$ and once weighted by the luminosity. The metallicity profiles were computed in spherical shells. The profile shows that the metallicity falls below 0.1 solar outside the disc radius of $\sim 8 \text{ kpc}$, and it drops to much smaller values outside the inner $\sim 30 \text{ kpc}$, where the streams are basically made of compressed intergalactic gas. Therefore dust opacity is estimated to reduce the Lyα luminosity from the streams by a small factor only. However, this assumption is likely to fail in the inner galactic disc, where the metallicity is above 0.1 solar despite the high redshifts. The effective transmission parameter $f_0$ in the galaxy vicinity is lower than that in the streams. By ignoring the dust, we overestimate the Lyα luminosity from the disc.

5 THE Lyα EMITTING GAS

Fig. 5 shows the cumulative mass-weighted temperature distribution in the CDB galaxies at $z = 2.3$. The distribution is specified alternatively for the halo gas in the radius range $(0.2-1.0) R_e$, and for the whole gas, including the inner part that involves the disc and its neighbourhood. We see that the temperature distribution in the halo is bimodal, with a hot virial phase at $10^5 < T < 2 \times 10^6 \text{ K}$ containing $\sim 35$ per cent of the gas and a cold phase (marked by a box) at $10^4 < T < 3 \times 10^3 \text{ K}$ containing $\sim 50$ per cent of the gas. A very cold tail at $T < 10^4 \text{ K}$ contains $\sim 15$ per cent of the gas.

For the somewhat more massive MN galaxies, which we do not show here, the situation is qualitatively similar. The temperature distribution is again bimodal, with a hot phase at $6 \times 10^4 < T < 5 \times 10^5 \text{ K}$ containing $\sim 45$ per cent of the gas and a cold phase at $2 \times 10^4 < T < 4 \times 10^3 \text{ K}$ containing $\sim 35$ per cent of the gas. An intermediate component in the range $4 \times 10^4 < T < 6 \times 10^4 \text{ K}$ contains $\sim 20$ per cent of the gas. In the two kinds of simulations, while the hot phase is spread throughout the halo, the cold phase is in the narrow streams flowing through the hot medium and the very cold tail is concentrated in the dense clumps within the streams and in the disc. Based on the temperature dependence of the emissivity in Fig. 3, we conclude that most of the Lyα luminosity is expected to come from the cold streams, with about half the total gas mass participating in efficient Lyα emission.

Fig. 6 shows the cumulative mass-weighted density distribution in the CDB galaxies at $z = 2.33$ in the same radius zones. If the multiphase medium is in pressure equilibrium, the rank order of the cells by density and by temperature are exactly opposite of one another, so a comparison of Figs 5 and 6 reveals that the Lyα emitting cold streams have densities in the range from below $0.01 \text{ cm}^{-3}$ to above $0.1 \text{ cm}^{-3}$ (marked by a box). The situation for the MN galaxies is comparable.

Fig. 7 shows the 2D gas distribution in the temperature–density plane. It complements Figs 5 and 6 with information on the different gas phases present in the simulations. We see that much of the gas is not in pressure equilibrium, $nT = \text{const}$. The top and bottom panels display the distribution of mass and Lyα luminosity, respectively. We see that the luminosity peaks at $T = (1-2) \times 10^4 \text{ K}$ over a broad range of densities from $n = 0.001 \text{ cm}^{-3}$ and up. We interpret the $T > 10^4 \text{ K} \text{ gas with densities above } n = 0.1 \text{ cm}^{-3} \text{ outside the central galaxy to be in clumps along the streams. It seems to be associated with a few per cent of the mass that contribute about 25 per cent of the luminosity. A similar fraction of the luminosity comes form gas in the temperature range } T = (2-7) \times 10^4 \text{ K}.

In order to evaluate the possible error introduced in the MN galaxies by practically ignoring the emissivity from cells with densities $n > 0.1 \text{ cm}^{-3}$, we computed the fraction of the luminosity coming from such cells in the CDB galaxies. We find that this fraction is typically only $\sim 15$ per cent in the halo at $r > 0.2 R_e$, and it could be as high as 40 per cent in the central disc. By adding 15 per cent to the MN emissivities everywhere we obtain a good match between the total luminosities at a given halo mass in the CDB and MN simulations (see Fig. 12).

6 SIMULATED Lyα BLOBS

Figs 8 and 9 show sample images of simulated galaxies, two CDB galaxies and two MN galaxies. The haloes are of virial mass $M_V \simeq 4 \times 10^{11}$ and $10^{12} \text{ M}_\odot$, respectively, and the corresponding redshifts are $z = 2.3$ and 2.5.

The top panels present the neutral hydrogen column density as computed in Section 4 assuming CIE. The middle panels are maps of Lyα rest-frame surface brightness $S$, namely a fraction $f_0$ of the emissivity integrated along the line of sight, as emitted per unit area at the galaxy. The bottom panels show images of ‘observed’ surface brightness $I$ as an observer would see it, per unit area in the galaxy and at the telescope. The rest-frame surface brightness is converted to observed surface brightness via

$$I = \frac{S}{4\pi(1+z)^2}.$$
In order to obtain realistic images, we applied a Gaussian point spread function (PSF) with a 0.6 arcsec full width at half-maximum (FWHM) to mimic atmospheric distortions in good seeing conditions, and assumed a pixel size of 0.2 arcsec at the telescope.

One can read from the middle maps the fraction of the luminosity that comes from each of the different parts of the halo. In particular, we consider separately (i) the contribution from the disc galaxy and its near vicinity inside a circle of radius $0.2R_e$ and (ii) the contribution from the halo at $0.2–1.0R_e$. We see in the figure that the rest-frame surface brightness in the inner galaxy reaches values higher than $S \sim 10^{40}$ erg s$^{-1}$ kpc$^{-2}$ but over a limited area of $\sim 10^2$ kpc$^2$, thus contributing $L \sim a$ few $\times 10^{42}$ erg s$^{-1}$. The typical surface brightnesses in the halo are lower by a factor of a few, but the emission regions are spread over an area that is larger by an order of magnitude, thus providing a comparable contribution to the total luminosity. Most of the luminosity comes from the low-density streams, and a non-negligible fraction is from clumps associated with the streams.

Recall that the emission from the disc may be partly absorbed by dust, which is ignored here, but we do not expect substantial dust absorption in the streams in the outer halo, where the metallicity is 0.01–0.1 solar. We can thus consider the halo contribution to be a safe lower limit to the overall Ly$\alpha$ luminosity, and take the luminosity as estimated from the whole system of disc plus halo as an upper limit.

In the bottom panels we see blobs with surface brightness above $10^{-18}$ erg s$^{-1}$ cm$^{-2}$ arcsec$^{-2}$ extending to $\sim 100$ kpc. The images are clumpy and irregular. Their shapes are non-circular; they tend to be elongated and asymmetric, showing finger-like extensions. These sample images resemble the observed images of LABs, e.g. those shown in fig. 8 of Matsuda et al. (2004). Fig. 10 compares ‘observed’ surface brightness maps of two simulated CDB galaxies (from the bottom panels of Figs 8 and 9) with the images of the two most luminous LABs observed by Matsuda et al. (2004). The qualitative similarity of the irregular, clumpy and elongated LAB morphologies and their sizes is encouraging.

7 Ly$\alpha$ Luminosity as a Function of Mass and Redshift

The variation of average total Ly$\alpha$ luminosity from within the virial sphere as a function of halo mass and redshift, as derived from the simulations, is shown in Fig. 11. The two panels refer to the luminosity from the whole halo, streams and central galaxy, as an upper limit, and to the streams at $r > 0.2R_e$ only as a lower limit. The fraction of the radiation emitted is set to $f_e = 0.85$, as determined by a fit to the observed luminosity function in Section 8. The luminosities are quoted for small bins about points in the $(M_V, z)$ plane. Four points refer to the average luminosities over the three CDB galaxies at $z = 1.9, 2.5, 3.2$ and 4.0. Each of the other points refers to the average over 12 MN galaxies. The 2D functional fit to the quoted values, $L(M_V, z)$, is shown in colours. We use a function of the form

$$L_{43}(M_V) = A(M_{12})^B(1 + z)^C,$$

where $A$, $B$ and $C$ are free parameters, $L_{43} \equiv L_{Ly\alpha}/10^{43}$ erg s$^{-1}$ and $M_{12} \equiv M_*/10^{12} M_\odot$. To determine the free parameters, we first address the mass dependence at a fixed redshift where we have performed our most detailed analysis in CDB, $z \simeq 2.4$. We then...
assume that the same mass dependence is valid between \( z \sim 2 \) and 4, and evaluate the redshift dependence.

Fig. 12 shows the luminosity as a function of halo mass at \( z = 3.1 \). The luminosity is computed within the isophotal surface brightness threshold of \( 2.2 \times 10^{-18} \) erg s\(^{-1}\) cm\(^{-2}\) arcsec\(^{-2}\), also used by Matsuda et al. (2004). We adopted \( f_\alpha = 0.85 \). A small correction upward of \( \sim 30 \) per cent has been applied to the results as extracted at \( z \sim 2.4 \) in order to bring them to \( z = 3.1 \) in this plot, using equation (21) with \( C = 1.3 \) as determined below. The average luminosity for the three CDB galaxies is shown at \( M_v \simeq 4 \times 10^{11} \) M\(_\odot\). Each of the other symbols is the average of 12 MN galaxies. We see that in the mass range most relevant for LABs, \( M_v = 10^{11.5} - 10^{12.5} \) M\(_\odot\), the luminosity is roughly proportional to halo mass. This dependence is driven by the near-linear mass dependence of the accretion rate, equation (5).

For a more accurate description of the mass dependence we fit the function of equation (21) to the results at the fixed redshift \( z \simeq 2.4 \), using least-squares fit in the log, and obtain the lines shown in Fig. 12 with \( B = 0.80 \) for the whole halo and \( B = 0.76 \) for the streams only.

We next use this functional mass dependence to scale the simulated results at other redshifts to a fixed mass, \( M_v = 10^{12} \) M\(_\odot\). The results are the values as quoted in the \((M_v, z)\) bins shown in Fig. 11. The scaled luminosities are shown in Fig. 13. A crude power-law fit to the symbols yields a redshift-dependent power index \( C = 1.3 \), with the normalization parameters \( A = 0.188 \) and 0.0972 for streams + galaxy and for streams only, respectively. We note that the obtained redshift dependence of the luminosity is somewhat weaker than the \((1+z)^2\) dependence of the accretion rate in equation (5). The accuracy of this scaling with redshift is not an important issue.
Figure 8. Images of two simulated CDB galaxies at $z = 2.3$ (left- and right-hand panels). The virial masses are $M_v \simeq 4 \times 10^{11} M_\odot$. The box side is 140 kpc (physical). The outer circle marks the virial radius and the inner circle is at 0.2$R_v$. Top: neutral hydrogen column density. Contours are shown for $10^{20}$ and $10^{21}$ cm$^{-2}$. Middle: rest-frame surface brightness $S$, with $f_\alpha = 0.85$, showing contours at $10^{39}$ and $10^{40}$ erg s$^{-1}$ kpc$^{-2}$. Bottom: ‘observed’ surface brightness $I$, at an angular resolution of $\sim 0.1$ arcsec, with contours at $10^{-18}$ and $10^{-17}$ erg s$^{-1}$ cm$^{-2}$ arcsec$^{-2}$, corresponding to the contours of $S$. The fraction of luminosity that originates from within these contours is 80 and 20 per cent, respectively.
Figure 9. Images of two simulated MN galaxies at $z = 2.5$ (left- and right-hand panels). The virial masses are $M_v \simeq 10^{12} M_\odot$. The box side is 184 kpc (physical). The outer circle marks the virial radius and the inner circle is at 0.2$R_v$. Top: neutral hydrogen column density. Contours are shown for $10^{20}$ and $10^{21}$ cm$^{-2}$. Middle: rest-frame surface brightness $S$, with $f_\alpha = 0.85$, showing contours at $10^{39}$ and $10^{40}$ erg s$^{-1}$ kpc$^{-2}$. Bottom: 'observed' surface brightness $I$, at an angular resolution of $\sim$0.1 arcsec, with contours at $10^{-18}$ and $10^{-17}$ erg s$^{-1}$ cm$^{-2}$ arcsec$^{-2}$, corresponding to the contours of $S$. The fraction of luminosity that originates from within these contours is 93 and 18 per cent, respectively.
Figure 10. A qualitative crude comparison of images of simulated (top) and observed (bottom) LABs. The observed LABs are the two most luminous LABs from Matsuda et al. (2004) with $L \simeq 10^{44}$ erg s$^{-1}$. The simulated images are of CDB galaxies scaled from $z = 2.3$ to 3.1 and to the luminosity of the corresponding observed LABs using $I \propto R \propto L^{1/3}$ (the original, unscaled luminosity was $5 \times 10^{43}$ erg s$^{-1}$). All four maps have an angular resolution of $\sim 0.2$ arcsec.

Figure 11. Total Ly$\alpha$ luminosity as a function of halo mass and redshift, separately for streams + galaxy (left-hand panel) and for streams only (right-hand panel). The average values as drawn from the simulations in bins of $(M_v, z)$ are marked by symbols and numbers in units of $10^{43}$ erg s$^{-1}$. The CDB results are marked by filled symbols and the MN results by open symbols. Shown in colour is the fitting function equation (21) with contours at $10^{42}$, $10^{43}$ and $10^{44}$ erg s$^{-1}$.
Cold streams as Lyα blobs

for us here since we only use it to correct the luminosities from \(z \simeq 2.4\) to 3.1, a correction of less than 30 per cent.

8 Lyα Luminosity Function

To obtain a predicted LAB luminosity function at \(z \sim 3\), we convolve \(L_{\text{Ly} \alpha}(M_v)\) with the Sheth–Tormen halo mass function (Sheth & Tormen 1999) at the same redshift. Here we assume a cosmology with \(\Omega_\Lambda = 0.72, \Omega_M = 0.28, \sigma_8 = 0.83, h = 0.71\) and \(\Omega_B = 0.045\). The resultant luminosity function is presented in Fig. 14, both for the total luminosity within the virial radius and for the halo streams only, in the range 0.2–1.0\(R_v\), which we consider to be upper and lower limits, respectively, given the uncertainty in the dust obscuration from the inner galaxy. The fraction of emitted radiation, which could range in principle between 0.5 and 1.0, is set to \(f_\alpha = 0.85\).

We computed a crude estimate of the observed Lyα luminosity function at \(z \sim 3.1\) in three luminosity bins using the most recent 201 LABs observed by Matsuda et al. (private communication). This sample contains three fields of different environment densities, and can thus be crudely considered as a fair sample. The number densities were determined by counting the number of blobs in the luminosity bins and dividing by the comoving survey volume of \(V_s = 1.5 \times 10^6\) Mpc\(^3\) (or physical volume 3.6 \(\times 10^5\) Mpc\(^3\)). Also shown is the luminosity function derived earlier from a partial sample of 35 LABs in a dense cluster environment (Matsuda et al.

![Figure 12](image1.png)

**Figure 12.** Lyα luminosity as a function of halo mass at \(z = 3.1\) with \(f_\alpha = 0.85\), for the total luminosity within the virial radius (upper, red symbols and curve), and for the halo outside 0.2\(R_v\), (lower, blue symbols and curve). Each symbol represents the average luminosity over simulated haloes of a given mass, 12 haloes from the MN simulation (open symbols) and three from the CDB simulations (filled symbols). The lines are least-squares fits using equation (21) with slopes \(B = 0.80\) and 0.76, respectively.

![Figure 13](image2.png)

**Figure 13.** Lyα luminosity as a function of redshift for \(M_v = 10^{12}\) M\(_\odot\) with \(f_\alpha = 0.85\), for the total luminosity within the virial radius (upper, red symbols and curve), and for the halo outside 0.2\(R_v\), (lower, blue symbols and curve). Each symbol represents the average luminosity over simulated haloes of a given mass, 12 haloes from the MN simulation (filled symbols) and three from the CDB simulations (open symbols). The luminosities were scaled from the simulated values at the different halo masses to \(M_v = 10^{12}\) M\(_\odot\) using the mass dependence obtained in equation (21). The curves are approximate eye-ball power-law fits with a slope \(C = 1.3\).

![Figure 14](image3.png)

**Figure 14.** Lyα luminosity function, showing the comoving number density of LABs brighter than \(L_{\text{Ly} \alpha}\) at \(z = 3.1\). The observed symbols (black, filled circles) are based on a sample of 201 LABs by Matsuda (private communication) in a fair sample. The earlier, higher estimates (magenta, open circles) are based on 35 LABs in a dense cluster environment. The lower and upper limits for the luminosity function as derived from the simulations are shown, for streams only (dashed, blue curve) and for streams + galaxy (solid, red curve), respectively. The transmission factor was set to \(f_\alpha = 0.85\) for a best match between theory and observation. The gravitational heating toy-model prediction based on equation (7) is shown (green, dashed line).

© 2010 The Authors. Journal compilation © 2010 RAS, MNRAS 407, 613–631
The latter overestimates the universal luminosity function by a factor of a few as expected from the high environment density. A Poissonian error has been attached to every bin. The number density has been derived by dividing the number of objects in each bin by the total survey volume.

We learn from Fig. 14 that the predicted luminosity function has a similar slope to the observed one. For $f_\alpha \sim 0.85$, the predicted upper and lower limits border the observed function from above and below throughout the whole range of LAB luminosities. We find that $f_\alpha$ changes roughly linearly with the adopted cosmological $\sigma_s$ through its effect on $m(M)$. Therefore, the $\sim 10$ per cent uncertainty in $\sigma_s$ translates into a similar uncertainty in $f_\alpha$.

Also shown is the luminosity function as estimated from our crude toy model for gravitational heating, Section 2, again convolved with the Sheth–Tormen halo mass function, and using $f_{df_c} = 0.34$. It is remarkable that this very crude toy model recovers the simulation results in the relevant mass range to within a factor of a few.

9 SURFACE BRIGHTNESS PROFILE, AREA AND LINENUMTH

A detailed comparison of theory to data will be presented in a future paper where we apply a more accurate radiative transfer calculation to the simulated galaxies. However, it is worthwhile to report here on a preliminary comparison with data, based on the simplified analysis of this paper, which is encouraging.

9.1 Surface brightness profile

Fig. 15 shows the surface brightness profile of the stacked images from the three CDB simulations, assuming $f_\alpha = 0.85$ throughout the whole halo. It is compared to stacked profiles of 35 observed LABs from Matsuda et al. (2004). For the stacking of the observed LABs, which range over more than an order of magnitude in luminosity, we scale the LAB radius $R$ and surface brightness $I$ to a fiducial luminosity $L_0 = 10^{18}$ erg s$^{-1}$ and assume $L \propto M_c$ and $R \propto R_c \propto M_c^{1/3}$ to obtain the scaling $R \propto I \propto L_0^{1/3}$. Thus, for a LAB with luminosity $L$, we multiply both the radius and the surface brightness by the factor $(L_0/L)^{1/3}$. Non-zero values for $I$ are considered for each LAB only above the isophotal surface brightness threshold of $2.2 \times 10^{-18}$ erg s$^{-1}$ cm$^{-2}$ arcsec$^{-2}$ and only the inner adjacent non-zero values (i.e. all values outside the innermost zero value are also treated as zero values) were used in the averaging of $I$ at each radius $r$. The stacked simulated profile has been scaled accordingly to match the same fiducial luminosity $L_0$.

We see that a power law $I(r) \propto r^{-\gamma}$ is a good fit to the simulated profile, with $\gamma \simeq 1.2$, from the disc scale $r < 10$ kpc out to $\sim 40$ kpc, which is about half the virial radius. It seems to steepen to $\gamma \sim 2$ at larger radii.

The observed images of Matsuda et al. (2004) are subject to a PSF of $\sim 1$ arcsec, which is responsible for the flattening of the observed profile at $r \leq 10$ kpc. Given the high background outside the surface brightness threshold, the meaningful part of the observed stacked profile is limited to the range 10–30 kpc. In this range, the power law with a slope $\gamma = 1.2$ provides a good fit to the observed profile as well. We also see that the crude toy model of Section 2 provides a profile that is consistent with the simulated profile to within a factor of 2, as seen in Fig. 1.

9.2 Isophotal area

Two of the observable global quantities for each LAB are the area encompassed by an isophotal contour of a given threshold surface brightness, and the corresponding luminosity. In order to compare to the data of Matsuda et al. (2004), we applied to our simulated galaxies, scaled to $z = 3.1$, a threshold of $2.2 \times 10^{-18}$ erg s$^{-1}$ cm$^{-2}$ arcsec$^{-2}$.

Fig. 16 displays the areas and luminosities of our simulated LABs in comparison with the 35 observed LABs. Each of the three simulated CDB galaxies is ‘observed’ from three orthogonal directions, and each of the randomly selected MN galaxies is ‘observed’ from one random direction. We see that the area at a given luminosity in MN is systematically smaller than in CDB, but only by a factor of 2. The simulations agree with the observations to within a factor of 2, and they reproduce the general correlation between area and luminosity.
The overall trend follows the scaling relation

$$A \propto L^{0.75}. \quad (22)$$

### 9.3 Linewidth

The major contribution to the LAB line profile is the kinematic effect owing to the distribution of line-of-sight velocities in the cold gas. It is dominated by the coherent instreaming velocities from the background and from the foreground, corresponding to a FWHM of the order of a few times the halo virial velocity. The effect of resonant scattering of Lyα should be convolved with the kinematic effect.

In Fig. 17 we show the kinematic components of three characteristic line profiles from the simulated CDB galaxies. The effect of resonant scattering is not included yet. For any desired direction of view, the kinematic line profile is computed as the luminosity-weighted distribution of the line-of-sight velocities across the whole isophotal area. In some cases, the line profile is dominated by one peak, corresponding to a stream that stretches roughly along the line of sight. In other cases, the line profile is bimodal, with a red peak and a blue peak corresponding to two opposite streams that lie not far from the line of sight. A third type of line profile is flatter, showing three peaks, the central of which corresponds to a stream that is roughly perpendicular to the line of sight. We crudely estimate the kinematic FWHM as twice the standard deviation of \( \Delta v \).

Resonant scattering of the Lyα photons in the stream neutral hydrogen gas broadens the Lyα emission line profile as the trapped photons diffuse into the damping wings before finally escaping (Adams 1972; Neufeld 1990). For scattering and escape from a static plane parallel medium, the total linewidth, in units of the luminosity, is

$$\sigma_{\text{scat}} \approx 2 \alpha \tau_{\text{Lyα}} \sqrt{(3/\pi)} 1^{1/3}. \quad (23)$$

For Lyα, the Doppler parameter is

$$b = 12.89 \ T_4^{1/2} \ \text{km s}^{-1},$$

and the radiation parameter \( \alpha \) and the optical depth \( \tau_{\text{Lyα}} \) are given in Section 4 following equation (19). We obtain

$$\sigma_{\text{scat}} \approx 436 \alpha_{20} T_4^{1/3} \ \text{km s}^{-1}. \quad (24)$$

Thus, the line profile resulting from resonant scattering of a mid-plane source is expected to consist of a red and a blue peak separated by \( \sigma_{\text{scat}} \). For a very crude estimate of total linewidth \( \Delta v \), we add in quadrature \( \sigma_{\text{scat}} \) to the FWHM of the kinematic line profile.\(^5\)

Fig. 18 displays isophotal area versus linewidth for the simulated LABs and for observed Lyα emitters Matsuda et al. (2006), including LABs and less extended emitters. The simulations and observations show a similar general trend, which very crudely follows \( A \propto (\Delta V)^{2.25} \), but the data show large scatter. A comparison to the tighter correlation seen in Fig. 16 between the area and the luminosity indicates that the scatter in Fig. 18 is mostly due to scatter in \( \Delta v \). This scatter is partly due to the variations in the relative orientations of the streams and the line of sight. The spread in the \( \Delta v \) as estimated in the simulations is underestimated due to the fact that we added in quadrature a constant value of \( \sigma_{\text{scat}} = 436 \ \text{km s}^{-1} \). This is also responsible for the artificially sharp lower bound at \( \Delta v = 436 \ \text{km s}^{-1} \), which is apparent for small values of \( A \).

Finally, Fig. 19 displays linewidth versus luminosity for the simulations and a subset of the observed LABs by Matsuda et al. (2006). The simulations reveal a scaling relation that can be approximated by \( \Delta v \propto L^{0.18} \). The simulations and observations agree to within a factor better than 2, but the larger scatter in the observed data

\(^5\) We note that the effect of resonant scattering as determined for a static medium may be an overestimate because of turbulent motions.
confirms our suspicion that we have underestimated the scatter in the analysis of $\Delta v$ from the simulations.

We conclude that our simplified analysis of $\text{Ly}\alpha$ emission from simulated galaxies qualitatively reproduces the observed correlations between the global quantities of $\text{Ly}\alpha$ luminosity, area and linewidth. A more detailed comparison of theory and observed LABs, especially involving line profiles and linewidths, should await a more accurate analysis of the simulations using radiative transfer and including dust absorption, as well as photometric and spectroscopic measurements of a larger sample of observed LABs.

## 10 Gravitational Heating

We now return to the energy source for the $\text{Ly}\alpha$ emission in our simulations. In addition to the gravitational heating, which was crudely approximated in Section 2, there is photoionization by the UV background. Photoionization effects by central sources such as stars and AGNs were not simulated, and they are not expected to be substantial because of the shielding implied by the radial orientation of the cold streams. Before we address gravitational heating again, we should verify the contribution of the UV background.

A simple estimate is as follows. For an isotropic background with mean photon intensity $4\pi J^{\ast}$ and mean photon energy $\epsilon$, the rate at which energy is absorbed across a spherical surface of radius $R$ is $\pi R^{2}J^{\ast}e4\pi R^{2}$. If we adopt for the metagalactic field at $z = 3$ a flux $\tau R^{\ast} \sim 5.5 \times 10^{4}$ photons s$^{-1}$ cm$^{-2}$ and $\epsilon \sim 20$ eV, we obtain a total UV heating rate into the virial radius $R_{v} \sim 100$ kpc of $2 \times 10^{43}$ erg s$^{-1}$. Even if all of the UV energy goes into heating the cold streams, the UV heating is small compared to the total $\text{Ly}\alpha$ luminosity of $\sim 10^{43}$ erg s$^{-1}$ and to the gravitational heating rate as estimated in Section 2. However, one should keep in mind that the MN simulation may overpredict the energy coming from the UV background when assuming that the gas is optically thin and the UV background is uniform.

To evaluate the relative contributions of gravitational heating and UV background in our actual simulation, we read from each grid cell of the simulation snapshot the instantaneous input rate of energy per unit volume by the UV flux, $\epsilon_{UV}$, and subtract it from the $\text{Ly}\alpha$ emissivity computed in equation (9), to obtain the relative contribution of gravitational heating in that cell,

$$f_{\text{grav}} \equiv \frac{\epsilon_{\text{Ly}\alpha} - \epsilon_{UV}}{\epsilon_{\text{Ly}\alpha}}.$$  \hspace{0.5cm} (25)

Fig. 20 shows the cumulative luminosity-weighted distribution of $f_{\text{grav}}$ for the three CDB galaxies at $z = 2.3$. In this plot $f_{\text{grav}} = 0$ refers to cells where all the $\text{Ly}\alpha$ luminosity is driven by the UV background and $f_{\text{grav}} = 1$ corresponds to gravitational heating only. The area under the curve is the fraction of the total energy provided by gravity. In the CDB haloes this fraction is 86.8 per cent, with only 13.2 per cent due to the UV background for the total halo within $R_{v}$. It is 84.2 and 15.8 per cent, respectively, for the halo between 0.2 and 1.0$R_{v}$. In more detail, if we focus, for example, on the cells of gas with the highest values of $f_{\text{grav}}$ that contribute 0.8 of the total luminosity, we read from the figure that they all have $f_{\text{grav}} > 0.8$. This means that in the gas that is responsible for 80 per cent of the $\text{Ly}\alpha$ luminosity, more than 80 per cent of the energy is gravitational and less than 20 per cent is due to the UV background. The fraction of the total energy provided by gravity in the MN haloes is 91.4 per cent, with only 8.6 per cent due to the UV background for the total halo within $R_{v}$. It is 79.5 and 20.5 per cent, respectively, for the halo between 0.2 and 1.0$R_{v}$. We conclude
that most of the Ly\(\alpha\) emission from our simulated galaxies is indeed driven by gravitational heating.

The total luminosity as predicted by the simplified gravitational heating toy model can be compared to the total Ly\(\alpha\) luminosity that is actually produced in the simulations. We focus on the halo streams in the radius range \(r = (0.2\text{--}1.0)R_\text{e}\) in haloes of \(M_\text{e} = 10^{12}\, M_\odot\) at \(z = 3\). The toy-model estimate, based on equation (7) and Fig. 1, is \(E_{\text{heat}} \sim 10^{43}\, \text{erg}\, \text{s}^{-1}\). A straightforward integral of the emissivity in the simulated galaxies over the same volume in the halo yields a comparable value (which can be read as \(L/f_\alpha\) from Fig. 12). This implies that the toy-model parameter \(f_\alpha\) should be of the order of unity, indicating that most of the gravitational energy is deposited in the cold streams.

We now return to Fig. 1, where we showed the gravitational heating toy-model prediction for the cumulative 3D luminosity profile. Shown in comparison is the corresponding Ly\(\alpha\) luminosity profile as derived from the actual simulated CDB galaxies, averaged over the three galaxies. The similarity of the two profiles is remarkable. This implies, again, that the Ly\(\alpha\) emission is indeed powered by the gravitational heating, where the potential energy of the instreaming cold gas is radiated as Ly\(\alpha\).

The cumulative luminosity profile can be fitted by a power law,
\[ L(<r) \propto r^\alpha, \]
(26)
corresponding to a 3D luminosity density profile \(\ell(r) \propto r^{\alpha-3}\). In the range \(r = (0.1\text{--}0.4)R_\text{e}\), we find for the simulated luminosity \(\alpha \simeq 0.8\), namely \(\ell \propto r^{-2.2}\). This is consistent with the average surface brightness profile \(I(r) \sim r^{-1.2}\) measured in Section 9.1. The toy model predicts a similar power index over the same radius range. We also note that the density profile of cold gas follows a similar power law.

\[ \frac{L(<r)}{L_{\text{total}}} \]

\begin{figure}[h]
\centering
\includegraphics[width=0.8\textwidth]{f20.png}
\caption{The role of gravitational heating versus photoionization. Shown is the cumulative, luminosity-weighted distribution of the fraction of the energy that is provided by gravitational heating, \(f_{\text{grav}}\), in the three CDB galaxies. The curves refer to the streams at \(r = (0.2\text{--}1.0)R_\text{e}\) (dashed blue) and to the whole halo including the inner galaxy (solid red). The area under the curve implies that for the total CDB halo within \(R_\text{e}\), 86.8 per cent of the energy input is gravitational, and only 13.2 per cent is due to the UV background. The numbers are 84.2 and 15.8 per cent, respectively, for the CDB halo between 0.2 and 1.0\(R_\text{e}\).}
\end{figure}

11 DISCUSSION AND CONCLUSION

Hydrodynamical cosmological simulations robustly demonstrate that massive galaxies at high redshifts were fed by cold gas streams, inflowing into dark matter haloes at high rates along the cosmic web (Keres et al. 2005; Dekel & Birnboim 2006; Dekel et al. 2009a).

In this paper we have shown that these streams should be observable as luminous Ly\(\alpha\) sources, with elongated irregular structures stretching for distances of over 100 kpc. The release of gravitational potential energy by the instreaming gas as it falls into the halo potential well is the origin of the Ly\(\alpha\) luminosity, which ranges from \(10^{42}\) to \(10^{44}\, \text{erg}\, \text{s}^{-1}\) for haloes in the mass range \(10^{13}\text{--}10^{15}\, M_\odot\) at \(z = 3\). The predicted Ly\(\alpha\) emission morphologies and luminosities make such streams likely candidates for the sources of observed high-redshift LABs. Most of the gas in the cold streams is at temperatures in the range \((1\text{--}5) \times 10^2\, \text{K}\), for which the Ly\(\alpha\) emissivity is maximized. The hydrogen gas densities in the streams are in the range 0.01--0.1 cm\(^{-3}\), and are higher in clumps that flow with the streams, leading to the high surface brightnesses.

Using state-of-the-art cosmological AMR simulations with 70-pc resolution and cooling to below \(10^3\, \text{K}\), and applying a straightforward analysis of Ly\(\alpha\) emissivity due to electron impact excitation, we produced maps of Ly\(\alpha\) emission from simulated massive galaxies at \(z = 3\). We computed the average Ly\(\alpha\) luminosity per given halo mass and the predicted luminosity function of these extended Ly\(\alpha\) sources, with an uncertainty of \(\pm 0.4\) dex. The properties of the individual images resemble those of the observed LABs in terms of morphology and kinematics. The predicted luminosity function is close to the observed LAB luminosity function. The simulated LABs qualitatively reproduce the observed correlations between the global LAB properties of Ly\(\alpha\) luminosity, isophotal area and linewidth.

The LAB properties can be understood using a very simple toy model that accounts for gravitational heating of the inflowing gas. In comparison, the more elaborate toy model of DL09 predicts a steeper power-law relation between luminosity and halo mass (their fig. 3, with \(f_\alpha = 0.2\)). In the mass range \(10^{12}\text{--}10^{13}\, M_\odot\), DL09 predict luminosities in the range \(1.5 \times 10^{42}\text{--}2.5 \times 10^{44}\, \text{erg}\, \text{s}^{-1}\). A comparison to Fig. 12 indicates that they underestimate the luminosity at \(M_\text{e} \sim 10^{13}\, M_\odot\) by a factor of a few. The DL09 toy model matches the observed luminosity function only after the predictions are boosted, trying to account for an overdense region of the universe. The association of the predicted LABs with massive haloes places them preferentially in overdense environments, as observed (Steidel et al. 2000; Matsuda et al. 2004, 2006), but this is already accounted for by the number density of haloes in a fair sample.

The association of observed LABs with star-forming LBGs (Hayashino et al. 2004), submillimetre galaxies (Chapman et al. 2001; Geach et al. 2005) or AGNs (Basu-Zych & Scharf 2004; Geach et al. 2009) is not surprising since star formation and AGN activity are triggered by the same process of streaming of cold gas into massive haloes. Ly\(\alpha\) emission can be driven by any of these sources of energy. For example, in photoionized gas in star-forming regions the Ly\(\alpha\) luminosity is \(L_{\text{Ly}\alpha} = 1.5 \times 10^{32}\, \text{SFR}\, \text{erg}\, \text{s}^{-1}\) where SFR is the SFR (in \(M_\odot\, \text{yr}^{-1}\)) for continuous star formation for a Kroupa IMF (Sternberg, Hoffmann & Pauldrach 2003). Thus the typical LAB luminosity of \(10^{43}\, \text{erg}\, \text{s}^{-1}\) would require SFR \(\sim 10\, M_\odot\, \text{yr}^{-1}\). However, most of this radiation is likely absorbed by dust, and confined to the central galaxy.

A limitation of our AMR simulations arises from the artificial pressure floor imposed in order to properly resolve the Jeans mass. This may have an effect on the temperature and density of the cold gas. The role of gravitational heating versus photoionization. Shown is the cumulative, luminosity-weighted distribution of the fraction of the energy that is provided by gravitational heating, \(f_{\text{grav}}\), in the three CDB galaxies. The curves refer to the streams at \(r = (0.2\text{--}1.0)R_\text{e}\) (dashed blue) and to the whole halo including the inner galaxy (solid red). The area under the curve implies that for the total CDB halo within \(R_\text{e}\), 86.8 per cent of the energy input is gravitational, and only 13.2 per cent is due to the UV background. The numbers are 84.2 and 15.8 per cent, respectively, for the CDB halo between 0.2 and 1.0\(R_\text{e}\).
gas in the streams, with potential implications on the computed Lyα emission. Still, the AMR code is the best available tool for recovering the stream properties. With 70-pc resolution and proper cooling below 10^4 K, the CDB simulations provide the most reliable description of the cold streams so far. The rather small correction that we had to apply to the luminosities extracted from the MN simulation indicates that the MN galaxies can be used to recover the mass and redshift dependence of the global stream properties and the scaling relations between them.

Another source of uncertainty has to do with the simplified way the ionization by the UV background is handled in the simulations, and with the post-processing calculation of the ionization state of the cold gas. In the CDB simulations, the centres of the streams, where the gas density is higher than 0.1 cm^{-3}, were assumed to be self-shielded against the UV background, in agreement with our analytic estimates in Section 4. The ionization state of the gas, which is a key ingredient in evaluating the Lyα emissivity, was computed in post-processing assuming CIE. Cooling rates were computed for the given gas density, temperature, metallicity and UV background based on CLOUDY (CDB) or assuming ionization equilibrium for H and He, including both collisional ionization and photoionization (MN). Photoionization is neglected since the streams are sufficiently thick to be self-shielded. An alternative calculation using CLOUDY yielded lower fractions of neutral hydrogen at n < 0.01 cm^{-3}, and total luminosities that are typically three times smaller than obtained using CIE. This calculation assumed that each cell, with a given hydrogen density n, is in the middle of a uniform slab of thickness 1 kpc. Based on our estimates of self-shielding, equation (17), we adopt the CIE results as the more reliable approximation.

In our computations so far we did not consider the radiative transfer of Lyα photons through the streams, or through the intervening IGM. We assume that most of the radiation emitted at z ~ 3 will reach the observer at z = 0 without undergoing significant attenuation. Resonant line scattering in the optically thick streams will tend to spread out the Lyα emission region, while the repeated Doppler shifts broaden the line profiles, and the increased path length of the random walk amplifies the absorption by dust. Such effects are expected to modify the images and line profiles, but to have only a small effect on the total Lyα luminosity. An analysis of Lyα emission from our simulated galaxies including the effects of radiative transfer and dust absorption and a more accurate treatment of photoionization from stars will be presented in a forthcoming paper (Kasen et al., in preparation).

Other hydrogen emission lines, such as Lyβ or Hα, are expected to be two orders of magnitude less luminous than Lyα in our model (Baldwin 1977; Miller 1974). The column densities of N_{H2} ~ 10^{20} cm^{-2} in the cold streams should also be detectable as Lyα absorption in quasar spectra (Prochaska 1999; Wolfe, Gawiser & Prochaska 2005). Considering all the galaxies that are intersected by a line of sight to a background quasar with an impact parameter <R_o, we crudely estimate from the simulations that absorption by N_{S2} > 1 hydrogen should be detected in ~20 per cent of the galaxies. Alternatively, Lyα photons that are emitted from the central galaxy should be absorbed in the radial streams feeding the same galaxy at N_{S2} > 10, but only in ~5 per cent of the galaxies. The streams could also be detectable as low-ionization metal absorbers (e.g. Gnat & Sternberg 2007) as long as the metallicity in the streams is greater than ~0.01 solar (paper in preparation).

Our results support the idea that the observed LABs are direct detections of the cold streams that drive the evolution of massive galaxies at high redshifts. Even though the observed LABs are sometimes associated with central sources that are energetic enough to power the observed Lyα emission, such as starbursts and AGNs, these central sources are very different from each other in the different galaxies, and in many LABs they are absent altogether (Geach et al. 2007; Yang et al. 2008). The gravitational heating associated with the inflowing cold streams is a natural mechanism for driving the extended Lyα cooling radiation observed as LABs, and this extended Lyα emission is inevitable in most high-redshift galaxies.

ACKNOWLEDGMENTS

The computer simulations were performed at NERSC, LBNL and the Barcelona Centro Nacional de Supercomputació as part of the Horizon collaboration. We thank Yuchi Matsuda, Kim Nilson and Masami Ouchi for sharing observational data with us. We acknowledge stimulating discussions with Nicolas Bouché, Michele Fumagalli, Orly Gnat, Daniel Kasen, Anatoly Klypin, Kamson Lai, Piero Madau, Ari Maller, Mark Mozena, Eyal Neistein, Hagai Netzer and Jason X. Prochaska. This research has been partly supported by an ISF grant, GIF I-895-207.7/2005, a France–Israel Teamwork in Sciences, the Einstein Center at HU and NASA ATP NAG5-8218 at UCSC. We thank the DFG for support via German–Israeli Project Cooperation grant STE1869/1-1.GE625/15-1. TG is a Minerva fellow and DC is a Gonda Meir fellow.

REFERENCES

Adams T. F., 1972, ApJ, 174, 439
Adams T. F., 1975, ApJ, 201, 439
Baldwin J. A., 1977, MNRA, 178, 67
Basu-Zych A., Scharf C., 2004, ApJ, 615, 85
Birnboim Y., Dekel A., 2003, MNRA, 345, 349
Birnboim Y., Dekel A., Neistein E., 2007, MNRA, 380, 339
Bullock J. S., Dekel A., Kolatt T. S., Kravtsov A. V., Klypin A. A., Porciani C., Primack J. R., 2001, MNRA, 321, 559
Callaway J., Unnikrishnan K., Oza D. H., 1987, Phys. Rev. A, 36, 2576
Ceverino D., Klypin A. A., 2009, ApJ, 695, 292
Ceverino D., Dekel A., Bournaud F., 2010, MNRA, 404, 2151 (CDB)
Ceverino-Rodríguez D., 2008, PhD thesis, New Mexico State Univ.
Chapman S. C., Lewis G. F., Scott D., Richards E., Borys C., Steidel C. C., Adelberger K. L., Shapley A. E., 2001, ApJ, 548, 17
Dalla Vecchia C., Schaye J., 2008, MNRA, 387, 1431
Dekel A., Birnboim Y., 2006, MNRA, 368, 2
Dekel A., Birnboim Y., 2008, MNRA, 383, 119
Dekel A. et al., 2009a, Nat, 457, 451
Dekel A., Sari R., Ceverino D., 2009b, ApJ, 703, 785
Dijkstra M., Loeb A., 2009, MNRAS, 400, 1109 (DL09)
Draine B., 2003, ARA&A, 41, 241
Dubois Y., Teyssier R., 2008, A&A, 477, 79
Elmegreen D. M., Elmegreen B. G., Ferguson T., Mullan B., 2007, ApJ, 663, 734
Fardal M. A., Katz N. G., Jeffrey P., Hernquist L., Weinberg D. H., Davé R., 2001, ApJ, 562, 605
Faucher-Giguère C., Prochaska J. X., Lidz A., Hernquist L., Zaldarriaga A., 2008, ApJ, 681, 831
Ferland G. J., Korista K. T., Verner D. A., Ferguson J. W., Kingdon J. B., Verner E. M., 1998, PASP, 110, 761
Förster Schreiber N. M. et al., 2006, ApJ, 645, 1062
Förster Schreiber N. M. et al., 2009, ApJ, 706, 1364
Furlanetto S. R., Schaye J., Springel V., Hernquist L., 2003, ApJ, 599, 1
Furlanetto S. R., Schaye J., Springel V., Hernquist L., 2005, ApJ, 622, 7
Geach J. et al., 2005, MNRA, 363, 1398
Geach J. et al., 2007, ApJ, 655, 9
Geach J. et al., 2009, ApJ, 700, 1
Genel S. et al., 2008, ApJ, 688, 789
