FIVE-YEAR WILKINSON MICROWAVE ANISOTROPY PROBE* OBSERVATIONS: ANGULAR POWER SPECTRA

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ABSTRACT

We present the temperature and polarization angular power spectra of the cosmic microwave background derived from the first five years of Wilkinson Microwave Anisotropy Probe data. The five-year temperature spectrum is cosmic variance limited up to multipole \( \ell = 530 \), and individual \( \ell \)-modes have signal-to-noise ratio \( S/N > 1 \) for \( \ell < 920 \). The best-fitting six-parameter \( \Lambda CDM \) model has a reduced \( \chi^2 \) for \( \ell = 33–1000 \) of \( \chi^2/\nu = 1.06 \), with a probability to exceed of 9.3%. There is now significantly improved data near the third peak which leads to improved cosmological constraints. The temperature-polarization correlation is seen with high significance. After accounting for foreground emission, the low-\( \ell \) reionization feature in the EE power spectrum is preferred by \( \Delta \chi^2 = 19.6 \) for optical depth \( \tau = 0.089 \) by the EE data alone, and is now largely cosmic variance limited for \( \ell = 2–6 \). There is no evidence for cosmic signal in the BB, TB, or EB spectra after accounting for foreground emission. We find that, when averaged over \( \ell = 2–6 \), \( \ell (\ell+1)c_{\ell}^{BB}/(2\pi) < 0.15 \mu K^2 \) (95% CL).

Key words: cosmic microwave background – cosmological parameters – cosmology: observations – early universe – large-scale structure of universe – space vehicles: instruments

1. INTRODUCTION

The Wilkinson Microwave Anisotropy Probe (WMAP) satellite (Bennett et al. 2003) has measured the temperature and polarization of the microwave sky at five frequencies from 23 to 94 GHz. Hinshaw et al. (2009) present our new, more sensitive temperature and polarization maps. After removing a model of the foreground emission from these maps (Gold et al. 2009), we obtain our best estimates of the temperature and polarization angular power spectra of the cosmic microwave background (CMB).

This paper presents our statistical analysis of these CMB temperature and polarization maps. Our basic analysis approach is similar to the approach described in our first-year WMAP temperature analysis (Hinshaw et al. 2003) and polarization analysis (Kogut et al. 2003) and in the three-year WMAP temperature (Hinshaw et al. 2007) and polarization analysis (Page et al. 2007). While most of the WMAP analysis pipeline has been unchanged from our three-year analysis, there have been a number of improvements that have reduced the systematic errors and increased the precision of the derived power spectra. Hinshaw et al. (2009) describe the WMAP data processing with an emphasis on these changes. Hill et al. (2009) present our more complete analysis of the WMAP beams based on five years of Jupiter data and physical optics fits to both the A- and B-side mirror distortions. The increase in main beam solid angle leads to a revision in the beam function that impacts our computed power spectrum by raising the overall amplitude for \( \ell > 200 \) by roughly 2%. Gold et al. (2009) introduce a new set of masks that are designed to remove regions of free–free emission that were a minor (but detectable) contaminant in analyses using the previous Kp2 mask used in the one- and three-year analysis. Wright et al. (2009) update the point source catalog presented by Hinshaw et al. (2007) finding 67 additional sources. Section 2 describes these changes and their implications for the measured power spectra.

Section 3 presents the temperature angular power spectrum (TT). WMAP has made a cosmic variance limited measurement of the angular power spectrum to \( \ell = 530 \) and we now report results into the “third peak” region. The WMAP results, combined with recent ground-based measurements of the TT angular power spectrum (Readhead et al. 2004; Jones et al. 2006; Reichardt et al. 2008), result in accurate measurements well into the “fifth peak” region. For WMAP, point sources are the largest astrophysical contaminant to the temperature power spectrum. We present estimates for the point source contamination based on multi-frequency data, source counts, and estimates from the bispectrum.

The polarization observations are decomposed into \( E \) and \( B \) mode components (Kamionkowski et al. 1997; Seljak &
Primordial scalar fluctuations generate only $E$ modes, while tensor fluctuations generate both $E$ and $B$ modes. With $T$, $E$, and $B$ maps, we compute the angular auto-power spectra of the three fields, $TT$, $EE$, and $BB$, and the angular cross-power spectra of these three fields, $TE$, $TB$, and $EB$. If the CMB fluctuations are Gaussian random fields, then these six angular power spectra encode all of the statistical information in the CMB. Unless there is a preferred sense of rotation in the universe, symmetry implies that the $TB$ and $EB$ power spectrum are zero. In Section 4 we present both the $TE$ and $TB$ temperature-polarization cross power spectra. The $WMAP$ measurements of the $TE$ spectrum now clearly see multiple peaks. The large angle $TE$ anti-correlation is a distinctive signature of superhorizon fluctuations (Spergel & Zaldarriaga 1997a). Komatsu et al. (2009) discuss how the $TB$ measurements constrain parity-violating interactions. Section 5 presents both the $EE$ and $BB$ polarization power spectra. The $EE$ power spectrum now shows a clear $\sim 5\sigma$ signature of cosmic reionization. Dunkley et al. (2009) show that the amplitude of the signal implies that the cosmic reionization was an extended process. Dunkley et al. (2009) and Komatsu et al. (2009) discuss the cosmological implications of the angular power spectrum measurements.

2. CHANGES IN THE FIVE-YEAR ANALYSIS

The methodology used for the five-year power spectra analysis is similar to that used for the three-year analysis. In this section we list the significant changes and their impact on the results:

1. Hinshaw et al. (2009) describe the changes in the map processing and the resultant reduction in the absolute calibration uncertainty from 0.5% to 0.2%.

2. The temperature mask used to compute the power spectrum has been updated, removing slightly more sky near the galactic plane, and more high-latitude point sources (Gold et al. 2009). The galactic mask used in the three- and one-year releases (Kp2) was constructed by selecting all pixels whose $K$-band emission exceeded a certain threshold. This procedure worked well in identifying areas contaminated by synchrotron emission; however, it missed a few small regions contaminated by free–free, particularly around $\rho$ Oph, the Gum nebula, and the Orion/Eridanus Bubble. For the five-year analysis we have constructed a new galactic mask to remove these contaminated areas.

Wright et al. (2009) updated the $WMAP$ point source catalog, finding 390 sources in the five-year data, 67 more sources than in the three-year catalog (Hinshaw et al. 2007). Of these 67 new sources, 32 were previously unmasked, and therefore added to the three-year source mask to create the five-year source mask.\footnote{Six of the sources in the five-year catalog were not added to the mask; they were found in a late update to the catalog after the mask had been finalized.}

All told, the new five-year temperature power spectrum mask (KQ85) retains 81.7% of the sky, while the three-year mask (Kp2) retained 84.6% of the sky.

3. The five-year polarization mask is the same as the three-year P06 mask described by Page et al. (2007), except that an additional 0.27% of the sky has been removed due to combining P06 with the new processing mask (Hinshaw et al. 2009).

4. In addition to masking, the maps are further cleaned of galactic foreground emission using external templates. The cleaning procedure is very similar to that of the three-year analysis; see Gold et al. (2009) for details. For the temperature map, three templates are used: a synchrotron template (the $WMAP$ $K - \alpha$ difference map), an Hz template as a proxy for free–free (Finkbeiner 2003), and a thermal dust template (Finkbeiner et al. 1999). For the polarization maps, two templates are used (since free–free is unpolarized): the polarized K-band map and a polarized dust template constructed from the unpolarized dust template, a simple model of the galactic magnetic field, and polarization directions deduced from starlight.

5. A great deal of work has gone into improving the determination of the angular power spectrum now clearly shows a clear $\sim 5\sigma$ signature of cosmic reionization. Dunkley et al. (2009) show that the amplitude of the signal implies that the cosmic reionization was an extended process. Dunkley et al. (2009) and Komatsu et al. (2009) discuss the cosmological implications of the angular power spectrum measurements.

3. TEMPERATURE SPECTRUM

The five-year $\ell \lesssim 32$ spectrum is described by Dunkley et al. (2009). At low-$\ell$ the likelihood function is no longer well approximated by a Gaussian so that we explicitly sample the likelihood function to evaluate the statistical distribution of each multipole.

We construct the five-year $TT$ spectrum for $\ell \gtrsim 32$ in the same fashion as the three-year spectrum; we refer the reader to Hinshaw et al. (2007) for details. For the five-year analysis we have constructed a new galactic mask to remove these contaminated areas.

Wright et al. (2009) updated the $WMAP$ point source catalog, finding 390 sources in the five-year data, 67 more sources than in the three-year catalog (Hinshaw et al. 2007). Of these 67 new sources, 32 were previously unmasked, and therefore added to the three-year source mask to create the five-year source mask.\footnote{12,582,912 pixels ($N_{\text{side}} = 1024$).}

All told, the new five-year temperature power spectrum mask (KQ85) retains 81.7% of the sky, while the three-year mask (Kp2) retained 84.6% of the sky.

3. The five-year polarization mask is the same as the three-year P06 mask described by Page et al. (2007), except that an additional 0.27% of the sky has been removed due to combining P06 with the new processing mask (Hinshaw et al. 2009).

We construct the five-year $TT$ spectrum for $\ell \lesssim 32$ in the same fashion as the three-year spectrum; see Gold et al. (2007) for details, and only briefly summarize the process as follows:

1. We start with the single year $V1$, $V2$, $W1$–$W4$ resolution-10 maps,\footnote{Six of the sources in the five-year catalog were not added to the mask; they were found in a late update to the catalog after the mask had been finalized.} masked by the KQ85 mask, and further cleaned via foreground template subtraction.

2. The pseudo-$C_\ell$ cross power spectra are computed for each pair of maps. Two weightings are used: flat weighting and inverse noise variance ($N_{\text{obs}}$) weighting.

3. The $\ell$-by-$\ell$ and $\ell$-by-$N_\ell$ auto-power spectra are not included in the combination, eliminating the need to subtract a noise bias.

4. A model of the unresolved point source contamination with amplitude $A_{\text{ps}} = 0.011 \pm 0.001 \mu K^2 sr$ is subtracted from the band-combined spectra. See Section 3.1 for more details.

5. The $V \times V$, $V \times W$, and $W \times W$ spectra are optimally combined $\ell$-by-$\ell$ to create the final CMB spectrum.

As in the three-year analysis, the diagonal elements of the $\hat{C}_\ell$ covariance matrix are calculated as

$$ (\Delta \hat{C}_\ell)^2 = \frac{2}{(2\ell + 1)J_{\text{sky}}^2(\ell)} (C_\ell + N_\ell)^2 $$

\footnote{Six of the sources in the five-year catalog were not added to the mask; they were found in a late update to the catalog after the mask had been finalized.}
where \( C_i \) is the cosmic variance term and \( N_l \) is the noise term. The value of \( f_{\text{sky}}(l) \), the effective sky fraction, is calibrated from simulations:\(^{17}\)

\[
f_{\text{sky}}(l) = \begin{cases} 
0.826 - 0.091(l/500)^2, & l \leq 500; \\
0.777 - 0.127(500/l), & l > 500. 
\end{cases} \tag{2}
\]

The five-year TT spectrum is shown in Figure 1. With the greater signal-to-noise ratio (S/N) of the five-year data the third acoustic peak is beginning to appear in the spectrum. The spectrum is cosmic variance limited up to \( \ell = 530 \) and individual \( \ell \)-modes have \( S/N > 1 \) for \( \ell < 920 \). In a fit to the best cosmological \( \Lambda \)CDM model, the reduced \( \chi^2 \) for \( \ell = 33-1000 \) is \( \chi^2/\nu = 1.06 \), with a probability to exceed of 9.3\%.

Figure 12 compares the unbinned five-year TT spectrum with the three-year result. Aside from the small upward shift of the five-year spectrum relative to that of the three-year, due to the new beam-transfer function, they are identical at low-\( \ell \). Figure 13 shows the unbinned TT spectrum broken down into its frequency components (\( V \times V, V \times W, W \times W \)), demonstrating that the signal is independent of frequency.

How much has the determination of the third acoustic peak improved with the five-year data? Over the range \( \ell = 680-900 \), which approximately spans the rise and fall of the third peak (from the bottom of the second trough to the point on the opposite side of the peak), the fiducial spectrum is preferred over a flat mean spectrum by \( \Delta \chi^2 = 7.6 \). For the three-year data it was \( \Delta \chi^2 = 3.6 \). With a few more years of data, WMAP should detect the curvature of third peak to greater than 3\( \sigma \).

In Figure 2 we compare the WMAP five-year TT power spectrum along with recent results from other experiments (Readhead et al. 2004; Jones et al. 2006; Reichardt et al. 2008), showing great consistency between the various measurements. Several on-going and future ground-based CMB experiments plan on calibrating themselves off their overlap with WMAP at the highest-\( \ell \)'s; improving WMAP’s determination of the third peak will have the added benefit of improving their calibrations.

### 3.1. Unresolved Point Source Correction

A population of point sources, Poisson-distributed over the sky, contributes an additional source of white noise to the measured TT power spectrum, \( C_{TT} \rightarrow C_{TT} + C_{ps} \). Given a known source distribution \( N(> S) \), the number of sources per steradian with flux greater than \( S \), the point-source-induced signal is

\[
C_{ps} = g(v) \int_0^S dS \frac{dN}{dS} S^2 \, (\mu K^2 \text{ sr}) \tag{3}
\]

where \( S \) is the source flux, \( S_{cl} \) is the flux cutoff (above which sources are masked and removed from the map), and \( g(v) = (e^v - 1)/2kv^2 \) converts flux density to thermodynamic temperature, with

\[
r(v) = \frac{(e^x - 1)^2}{x^2e^x}, \quad x \equiv hv/kT_{\text{CMB}} \tag{4}
\]

converting antenna to thermodynamic temperature.

At the frequencies and flux densities relevant for WMAP, source counts are dominated by flat-spectrum radio sources, which have flux spectra that are nearly constant with frequency (\( S \sim v^0 \) with \( \alpha \approx 0 \)). Wright et al. (2009) find the average spectral index of sources bright enough to be detected in the WMAP five-year data to be \( \langle \alpha \rangle = -0.09 \), with an intrinsic dispersion of \( \sigma_{\alpha} = 0.176 \). Since a source with flux \( S \sim v^{\alpha} \) has a thermodynamic temperature \( T \sim v^{\alpha-2}r(v) \), we model the frequency dependence of \( C_{ps} \) as

\[
C_{ps}(v_i, v_j) = A_{ps}(v_i)r(v_j) \left( \frac{v_i}{v_Q} \right)^{\alpha-2} \tag{5}
\]

where \( v_i, j \) are the frequencies of the two maps used to calculate the TT spectrum, \( A_{ps} \) is an unknown amplitude, and \( v_Q = 40.7 \text{ GHz} \) is the Q-band central frequency.

In this section, we estimate the value of \( A_{ps} \) needed to correct the TT power spectrum, finding \( A_{ps} = 0.011 \pm 0.001 \mu K^2 \text{ sr}, \)

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17 The Markov chains in Dunkley et al. (2009) and Komatsu et al. (2009) were run with a version of the WMAP likelihood code with older and slightly larger values for \( f_{\text{sky}} \). The change in \( f_{\text{sky}} \) increased the TT errors by on average 2\%. Retraining the \( \Lambda \)CDM chain with the new \( f_{\text{sky}} \) leads to parameter shifts of at most 0.1\( \sigma \).
and discuss incorporating its uncertainty into the likelihood function.

3.1.1. Estimating the Correction

For a fixed beam size, flat-spectrum radio sources are much fainter in the W-band temperature maps than in Q- or V-band, allowing us to use the frequency dependence of the TT spectrum at high-$\ell$ to constrain the value of $A_{ps}$. As in previous releases, the estimator we use is

$$\hat{A}_{ps} = \frac{\sum_{\alpha\beta} C^r_{\ell} (\Sigma^{-1})_{\alpha\beta} h^\beta_{\ell}}{\sum_{\alpha\beta} s^\alpha_{\ell} (\Sigma^{-1})_{\alpha\beta} h^\beta_{\ell}},$$

(6)

$$h^\gamma_{\ell} = s^{\gamma}_{\ell} - \frac{\sum_{\alpha\beta} s^\alpha_{\ell} (\Sigma^{-1})_{\alpha\beta} h^\beta_{\ell}}{\sum_{\alpha\beta} (\Sigma^{-1})_{\alpha\beta}},$$

(7)

where Greek letters represent a pair of frequencies (e.g., VW), $C^r_{\ell}$ is the measured TT cross-power spectrum, $\Sigma^{-1}$ is the $(C^r_{\ell} C^r_{\ell})^{-1}$ covariance matrix including cosmic variance and detector noise, and $s^\alpha_{\ell} = \ell (l+1) C^{\alpha\alpha}(\ell)/2\pi$. The inverse estimator variance $(\delta_{\hat{A}_{ps}})^{-2}$ is given by the denominator of (6). While $\Sigma^{-1}$ does not include the off-diagonal coupling due to the mask, the diagonal elements are renormalized to account for the loss of sky coverage.

Measured values for $A_{ps}$ are listed in Table 1 for various frequency combinations (QVW and VW) and galactic masks (KQ85, KQ80, and KQ75). The QVW estimates are insensitive to the galactic mask; the VW estimate increases somewhat as $\Delta$TQ increases. Both the QVW and VW estimates prefer the same value (KQ85, KQ80, and KQ75). The QVW estimates are insensitive to the power spectrum and $A_{ps}$ scales roughly as $\sim \nu^{-2}$ in the WMAP data. The red data points are from the analysis of V and W bands alone and the blue points are from the analysis of Q, V, and W bands. The horizontal dashed green lines, at 0.010 and 0.012, show the 1σ bounds for our adopted value of $A_{ps}$. Note that the QVW amplitude is independent of $\ell$.

| Bands | Mask | $A_{ps}(\ell = 0)$ (10^{-3} \mu K^2 sr) | $A_{ps}(\ell = -0.09)$ (10^{-3} \mu K^2 sr) |
|-------|------|---------------------------------|---------------------------------|
| QVW   | KQ85 | 11.3 ± 0.9                      | 11.2 ± 0.9                      |
|       | KQ80 | 11.3 ± 0.9                      | 11.2 ± 0.9                      |
|       | KQ75 | 10.7 ± 1.0                      | 10.6 ± 1.0                      |
| VW    | KQ85 | 6.9 ± 3.4                       | 7.2 ± 3.5                       |
|       | KQ80 | 9.1 ± 3.6                       | 9.5 ± 3.8                       |
|       | KQ75 | 10.5 ± 3.9                      | 11.1 ± 4.1                      |

Note. All results are for $\ell = 300$–800.

Figure 3. Unresolved point source contamination $A_{ps}$, measured in bins of $\Delta \ell = 100$ evaluated at 40.7 GHz (Q band). For a source population whose fluxes are independent of frequency $A_{ps}$ scales roughly as $\sim \nu^{-2}$ in the WMAP data. The red data points are from the analysis of V and W bands alone and the blue points are from the analysis of Q, V, and W bands. The horizontal dashed green lines, at 0.010 and 0.012, show the 1σ bounds for our adopted value of $A_{ps}$.

Figure 4. TT V–W null spectrum. After correcting for unresolved point source contamination, the individual power spectra are subtracted in power spectrum space. The result is consistent with zero and thus there is no evidence of point source contamination. In these units, point source contamination would be evident as a horizontal offset from zero. At $\ell = 500$, the TT power spectrum is $C_{\ell}^{TT} = 0.06$; thus the contamination is limited to roughly 3% in power.

et al. (1998, Tof98)18 and de Zotti et al. (2005, deZ05). $C^{ps}$ is calculated via (3), and $b^{ps}$ from

$$b^{ps} = 3 g(v) \int_{0}^{S_c} dS \frac{dN}{dS} S^{3},$$

(8)

where $g(v)$ and $S_c$ are defined in (3). The comparison is complicated by the fact that $S_c$ is unknown. We mask out not only the sources detected in WMAP data, but also undetected sources from external catalogues that are likely to contribute contaminating flux. However, a single value of $S_c$ predicts both $C^{ps}$ and $b^{ps}$, so we can in principle tune $S_c$ to match one, and see if it agrees with the other. In Table 2 we compare our measured values of $C^{ps}$ and $b^{ps}$ with the rescaled To98 and deZ05 predictions for several values of $S_c$. There is some tension

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18 In Bennett et al. (2003) we found that the To98 model needed to be rescaled by a factor of 0.66 to match the WMAP one-year number counts; Wright et al. (2009) refined the rescaling factor to 0.64 to match the WMAP five-year source counts.
between the measured values and the model predictions. Given our measured value for $b_{ps}$ the models would prefer a smaller value for $A_{ps}$, in the range 0.008–0.010 $\mu K^2$ sr. For the Tof98 model, the $S_{\delta}$ ≈ 0.52 predictions are within 1σ of both $C_{\ell}^\delta$ and $b_{ps}$. However, the deZ05 model appears to be discrepant, and a single value for $S_{\delta}$ cannot match both $C_{\ell}^\delta$ and $b_{ps}$.

Other groups have independently estimated the unresolved source contamination, and their results are in general agreement with ours. When the three-year data were initially released the correction was $A_{ps} = 0.017 \pm 0.002 $ $\mu K^2$ sr. Huffenberger et al. (2006) reanalyzed the data and claimed $A_{ps} = 0.011 \pm 0.001$, noticing that $A_{ps}$ was sensitive to the choice of galaxy mask: using the KP0 mask instead of KP2 reduced the value of $A_{ps}$. Revisiting our original estimate for the three-year analysis, we reduced the correction to 0.014 ± 0.003 for the published papers. In a subsequent paper, Huffenberger et al. (2008), the same group corrected their original estimate after finding a small error, finding 0.013 ± 0.001, consistent with our published result.

### 4. TEMPERATURE-POLARIZATION SPECTRA

The standard model of adiabatic primordial density fluctuations predicts a correlation between the temperature and polarization fluctuations. The temperature traces primarily the density, and $E$-mode polarization the velocity, of the photon–baryon plasma at recombination. The correlation was seen in earlier WMAP data by Kogut et al. (2003) and Page et al. (2007). The anti-correlation near $\ell = 30$ provides evidence that fluctuations exist on superhorizon scales, as is observed on an angular scale larger than the acoustic horizon at decoupling (Spergel & Zaldarriaga 1997b).

No significant changes have been made in the five-year TE analysis. We continue to use the method described by Page et al. (2007) to compute the TE power spectrum. The inputs are the KaQV polarization maps (Gold et al. 2009) and the VW temperature maps. For high multipoles $\ell > 23$, the likelihood can be approximated as a Gaussian, and we continue to use the ansatz given in Appendix C of Page et al. (2007) to compute the covariance matrix. At low multipoles, $\ell \lesssim 23$, the likelihood of the polarization data is evaluated directly from the maps, following Appendix D in Page et al. (2007).

Figure 5 shows the TE spectrum. At low-$\ell$ the spectrum and error bars are approximated using the Gaussian form, although these are not used for cosmological analysis. With five years of data the anti-correlation at $\ell = 140$ is clearly seen in the data, and the correlation at $\ell = 300$ is measured with higher accuracy. The second anti-correlation at $\ell \sim 450$ is now better characterized, and is consistent with predictions of the $\Lambda$CDM model. The structure tests the consistency of the simple model, which fits both the TT and TE spectra with only six parameters. The best-fit $\Lambda$CDM model has $\chi^2 = 415$ for the TE component, with 421 degrees of freedom, giving $\chi^2/\nu = 0.99$. The consistency confirms that the fluctuations are predominantly adiabatic, and constrains the amplitude of isocurvature modes.

The signal at the lowest multipoles, evaluated using the exact likelihood, is used to provide additional constraints on the reionization history. Although small, the measurement is consistent with the EE signal, and consistent with the three-year WMAP observations of Page et al. (2007). Dunkley et al. (2009) discuss constraints on reionization.

No correlation is expected between the temperature and the $B$-mode polarization. The TB spectrum is therefore primarily used as a null test, and is shown in Figure 6. It is consistent with no signal, as expected; over $\ell = 24–450$ the reduced null $\chi^2$ is 0.97. This measurement is used in Komatsu et al. (2009) to place constraints on the presence of any parity violating terms coupled to photons, that could produce a TB correlation. We now include the TB spectrum at high-$\ell$ as an optional module for the likelihood code.

### Table 2: Unresolved Point Source Contamination

| $S_{\delta}$ (Jy) | $C_{\ell}^\delta \times 10^{-3} \mu K^2$ sr | $b_{ps} \times 10^{-3} \mu K^2$ sr² |
|------------------|---------------------------------|---------------------------------|
| WMAP5 (KQ75)     | $\cdots$                         | 11.7 ± 1.1                      |
|                  |                                 | 4.3 ± 1.3                      |
| Toffolatti et al. (1998) $\times 0.64$ | 0.6                             | 12.1                           |
|                  |                                 | 6.8                             |
|                  |                                 | 0.5                             | 10.4                           |
|                  |                                 | 4.9                             |
| de Zotti et al. (2005) | 0.7                             | 11.7                           |
|                  |                                 | 8.4                             |
|                  |                                 | 0.5                             | 8.3                            |
|                  |                                 | 4.3                             |

Notes. All numbers are evaluated at 40.7 GHz (Q band).

a By Equation (5), $C_{\ell}^m(Q) = A_{ps}(Q)^2 = 1.089 \times A_{ps}$, where $A_{ps}$ is the QVW/KQ75 result from Table 1.

b From Komatsu et al. (2009), using the Q-band map and KQ75 mask.
Figure 7. Conditional likelihoods for the $\ell = 2$–7 EE multipole moments (black curves), computed using the WMAP likelihood code by varying the multipole in question, with all other multipoles fixed to their fiducial values. For example, in the $\ell = 4$ panel, the black curve is $f(x) \propto L(d|...,C_{EE}^4, C_{EE}^5, ...)$. For comparison, naïve pseudo-$C_\ell$ estimates are also shown with Gaussian errors (red curves). The pseudo-$C_\ell$ errors are noise only, while the conditional distributions include cosmic variance.

Table 3

| Likelihood Treatment | $n_s$         | $\sigma_8$       |
|----------------------|---------------|------------------|
| Standard             | $0.964 \pm 0.014$ | $0.796 \pm 0.036$ |
| SRCMARG              | $0.964 \pm 0.014$ | $0.798 \pm 0.036$ |
| SRCMARG $\times 5$  | $0.965 \pm 0.015$ | $0.803 \pm 0.042$ |
| BEAMMARG             | $0.964 \pm 0.015$ | $0.799 \pm 0.036$ |
| BEAMMARG $\times 20$| $0.958 \pm 0.016$ | $0.796 \pm 0.034$ |

Notes. One-dimensional marginalized values for $n_s$ and $\sigma_8$ for various treatments of the unresolved point source and beam uncertainty in the WMAP likelihood code. See Appendix A for descriptions of SRCMARG and BEAMMARG. Here, “$\times 5$” and “$\times 20$” indicate the error has been increased by a factor of 5 and 20, respectively.

5. POLARIZATION SPECTRA

Due to its thermal stability (Jarosik et al. 2007) and well-characterized gain, WMAP can measure polarization signals even though the scan pattern was not optimized for doing so. The polarization signal is manifested in the time-ordered data (TOD) differently from the temperature signal. As a result, some of the low-$\ell$ polarization multipoles are well sampled and other multipoles are poorly sampled and have large statistical errors (Hinshaw et al. 2007; Page et al. 2007). This is a rather different situation than from that of the temperature spectrum, and the data must be analyzed with some care.

When we analyze the $\ell = 2$ temperature power spectrum, we use the likelihood function rather than Gaussian errors, as the Gaussian approximation starts to break down with only $\approx 4$ effective modes measured in the map (the reduction is due to $f_{\text{sky}} \approx 0.7$). For polarization, this effect is even more dramatic, as our scan pattern significantly lowers the effective number of multipoles measured, particularly for EE $\ell = 2, 5, 7,$ and 9 and BB $\ell = 3$ (the peaks seen in Figure 16 in Page et al. 2007). Figure 8 demonstrates the importance of using the full likelihood description. The figure shows both the pseudo-$C_\ell$ estimates of the $\ell = 2$–7 BB multipoles and the conditional likelihoods computed using the WMAP likelihood code by varying the multipole in question, keeping the rest of the spectrum fixed to the fiducial best-fit $\Lambda$CDM model. From the plots it is clear...
Figure 8. Conditional likelihoods for the $\ell = 2$–7 BB multipole moments (black curves), computed using the WMAP likelihood code by varying the multipole in question, with all other multipoles fixed to their fiducial values. For comparison, naïve pseudo-$C_\ell$ estimates are also shown with Gaussian errors (red curves). The pseudo-$C_\ell$ errors are noise only, while the conditional distributions include cosmic variance. Note the large difference between the likelihood code and the pseudo-$C_\ell$ value for $\ell = 3$; this mode is sensitive to the time-ordered data baseline and is extremely poorly measured by WMAP, illustrating the complicated noise structure of the polarization data on large scales.

that the best estimates of the mean and the uncertainty are not attained with the pseudo-$C_\ell$ estimates.

We next consider the low-$\ell$ EE and BB power spectra in more detail. The low-$\ell$ EE power spectrum is shown in Figure 9. The uncertainties are obtained from the conditional likelihood and include cosmic variance; thus one cannot double the error flags to get the 95% confidence limits. If we zero out the $\ell < 10$ portion of the fiducial EE and TE spectra the $\chi^2$ increases by 22.3, of which 2.7 is due to TE. Thus, the reionization feature in the EE power spectrum is preferred by $\Delta \chi^2 = 19.6$. The $\ell = 2$, 3, 4, and 6 multipoles are cosmic variance limited, and the $S/N$ for the combined $\ell = 2$–7 bandpower is 11.

Considerable effort has gone into understanding the W-band $\ell = 7$ EE signal. Because of the apparently anomalously high-$\ell = 7$ EE value computed by the pseudo-$C_\ell$ algorithm, we have avoided using the W-band maps in cosmological analysis and use them only as an additional check on various models. Figure 8 of Hinshaw et al. (2009) shows that the $\ell = 7$ value, while high, appears to be consistent with being in the tail of a properly computed likelihood distribution. The W-band $\ell = 7$ problem may be a signature of poor statistics rather than a systematic. However, more data are needed to understand this potential anomaly. The $\ell = 3$ BB signal gives perhaps the clearest example of the importance of using the full likelihood code. While the pseudo-$C_\ell$ estimate implies a significant detection of power, the full likelihood code shows this to not be the case. The physical cause of the large uncertainty is that with our scan strategy an $\ell = 3$ BB signal resembles an offset in the data and thus is not well separated from the baseline (Page et al. 2007; Hinshaw et al. 2009).

We see no evidence for a B-mode signal at low-$\ell$, limiting the possible level to $\ell(\ell+1)C_{\ell \to 2-6}/(2\pi) < 0.15 \mu K^2$ (95% CL), including cosmic variance. With $\tau = 0.1$ and $r = 0.2$, a typical estimate for currently favored models of inflation, $\ell(\ell+1)C_{\ell \to 2-6}/(2\pi) \approx 0.008 \mu K^2$. Since a signal of 0.15 $\mu K^2$ corresponds roughly to $r \approx 20$, one can see that WMAP’s limit
is not based on the BB data, but on the tensor contribution to the TT and EE spectra as discussed by Komatsu et al. (2009).

For EE at $\ell > 10$, there are hints of signal in the data consistent with the standard ΛCDM model. However, the significance is not great enough to contribute to knowledge of the cosmological parameters. The five-year high-$\ell$ EE spectrum is shown in Figure 10, along with recent results from ground-based experiments (Leitch et al. 2005; Montroy et al. 2006; Sievers et al. 2007). For $\ell = 50$–800, $\chi^2 = 859.1$ assuming $C_{EE} = 0$, and drops by 8.4, or almost 3σ, assuming the standard ΛCDM model. For the three-year data the equivalent change in $\chi^2$ was 6.2.

The high-$\ell$ BB spectrum is consistent with no signal, having a reduced $\chi^2$ of 1.02 over $\ell = 50$–800 for the QV data. The lack of any signal in the low- and high-$\ell$ BB data is a necessary check of the foreground subtraction. As seen in Page et al. (2007), foreground emission produces $E$-modes and $B$-modes at similar levels; thus the absence of a $B$-mode signal suggests that the level of contamination in the $E$-mode signal is low. This is quantified by Dunkley et al. (2009).

6. SUMMARY AND CONCLUSIONS

We have presented the temperature and polarization angular power spectra of the CMB derived from the first five years of WMAP data. With greater integration time our determination of the third acoustic peak in the TT spectrum has improved. The low-$\ell$ reionization feature in the EE spectrum is now detected at nearly 5σ. The TB, EB, and BB spectra show no evidence for cosmological signals. The spectra are in excellent agreement with the best-fit ΛCDM model. Our knowledge of the power spectrum is improving both due to more detailed analyses, better modeling and understanding of the foreground emission, and more integration time.

All of the five-year WMAP data products are being made available through the Legacy Archive for Microwave Background Data Analysis (LAMBDA19), NASA’s CMB Thematic Data Center. The temperature and polarization angular power spectra presented here are available, as is the WMAP likelihood code which incorporates our estimates of the Fisher matrix, point sources, and beam uncertainties.

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APPENDIX

LIKELIHOOD TREATMENT OF SOURCE/BEAM UNCERTAINTIES

In this appendix, we test the treatment of the unresolved point source correction and beam uncertainties in the WMAP likelihood code, and show that it produces the correct results for cosmological parameters.

We adopt the same likelihood treatment of the unresolved point source correction uncertainty for the five-year likelihood code as used in the three-year code (Hinshaw et al. 2007, Appendix A), updated for the five-year value of $A_{ps}$. Briefly, a correction to the logarithmic likelihood, $L \equiv -2 \ln L = C_0 + L_1$, where $C_0$ is the standard likelihood and $L_1$ is the combined source and beam correction, is calculated assuming the $C_l$ are normally distributed, a reasonable assumption at high-$\ell$.

Huffenberger et al. (2007, Hut08) disagreed with the source and beam likelihood module used in the three-year analysis, pointing out that the uncertainty in $n_s$ (the index of primordial scalar perturbations) was unchanged even if the uncertainty in $A_{ps}$ was increased by a factor of 100 (Figure 2 in their paper). They proposed an alternative approach, integrating the beam/source covariance matrix into the cosmic variance/noise/mask covariance matrix and inverting the result in order to compute $L$ directly, instead of calculating $L_1$ as a separate correction. Using this form of the likelihood, as $\delta A_{ps}$ was increased, the uncertainty in $n_s$ increases (albeit modestly; $\delta n_s$ increased by 38% when $\delta A_{ps} \rightarrow 100 \times \delta A_{ps}$).

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19 http://lambda.gsfc.nasa.gov/
However, while we agree that it is striking that the error in $n_s$ is seemingly unaffected by the uncertainty in $A_{ps}$, we have some concerns regarding the Huf08 approach. To quote Huf08, “the errors on the source measurement do not make much difference, as long as $|\delta A_{ps}| < 0.003 [\mu K^2 \text{sr}]$,” and their Figure 2 implies the same holds true when $\delta A_{ps} = 0.003$. This value is significant because it is the uncertainty adopted for the three-year WMAP analysis. When Huf08 adopted the same uncertainty, they found the same absolute uncertainty in $n_s$ as the WMAP team, but their central value was shifted higher by 0.005. This shift persisted as $\delta A_{ps} \rightarrow 0$, and thus was seemingly not due to the point source uncertainty. The conclusion we draw is that they found the same result of $\delta n_s = 0.005$. This shift persisted as long as $\delta A_{ps} \rightarrow 0$, and thus was seemingly not due to the point source uncertainty. The conclusion we draw is that they found the same result of $\delta n_s = 0.005$. This shift persisted as long as $\delta A_{ps} \rightarrow 0$, and thus was seemingly not due to the point source uncertainty.

Huf08 quoted the value of $\ell_1$ computed with their alternative likelihood module for a particular CMB spectrum distributed with the WMAP three-year likelihood code test program, finding $\ell_1 = -2.64$, whereas the WMAP value is $\ell_1 = -1.22$. As a check, we numerically marginalize the $\ell_0$ portion of the likelihood over beam and point source errors, to see if we can reproduce their value. The desired integral is

$$\exp(-\ell_1/2) = \frac{1}{\mathcal{L}_0 d(C_{\ell})} \int dx d\vec{y} e^{-(x^2+y^2)/2} \mathcal{L}_0(d|C_{\ell}(x, \vec{y})),$$

where

$$C_{\ell}(x, \vec{y}) \equiv \left( C_{\ell}^{TT} + x \sigma_i^{\text{prec}} \right) \left( 1 + \sum_j y_j \sigma_{i,j}^{\text{beam}}(l) \right)$$

is the theoretical model ($C_{\ell}^{TT}$) perturbed by point source and beam errors. With 10 dimensions to integrate over (nine beam modes and one point source mode), normal grid-based quadrature is impractical, so we turn to Monte Carlo integration instead:

$$\exp(-\ell_{1,\text{MC}}/2) \approx \frac{1}{N_{\text{MC}}} \sum_{i=1}^{N_{\text{MC}}} \ln L(d|C_{\ell}(x^{(i)}, \vec{y}^{(i)})) - \ln L(d|C_{\ell}),$$

where $x^{(i)}$ and $y^{(i)}$ are independent unit-variance normal deviates. With $N_{\text{MC}} = 10^4$ points, we find $\ell_{1,\text{MC}} = -1.29 \pm 0.04$, consistent with the WMAP result of $-1.22$, but not the Huf08 result of $-2.64$.

As a further test of whether our cosmological parameter estimates fully capture the point source uncertainty, we have run a Markov chain with a modified form of the point source likelihood module, dubbed SRCMARG. The point source correction is calculated via a simple numerical integration,

$$\exp(-\ell_{1,\text{prec}}/2) = \int d\alpha \frac{1}{\sqrt{2\pi}} e^{-\alpha^2/2} \mathcal{L}_0(d|C_{\ell} + i \Delta \sigma_i^{\text{prec}})$$

$$\approx \frac{\Delta}{\sqrt{2\pi}} \sum_{i=-N}^{N} w_i e^{-i\Delta^2/2} \mathcal{L}_0(d|C_{\ell} + i \Delta \sigma_i^{\text{prec}})$$

with $N = 25$, $\Delta = 0.2$, and $w_i = 1$ except at the endpoints where $w_{|N|} = 1/2$ (the trapezoidal rule). The resulting one-dimensional marginalized distribution for $\sigma_8$, shown in the left panel of Figure 11, is indistinguishable from our standard result.
We have also run a SRCMARG chain with the error increased by a factor of 5 (i.e., \(\delta A_{\text{ps}} = 0.005\)). In this case the uncertainty in \(\sigma_8\) increases by 15%.

Likewise, we have run similar tests of the beam uncertainty, dubbed BEAMMARG. The approach is the same as SRCMARG, but with “\(C_l + \alpha \sigma^2_l,\text{beam}\) in (12) replaced by “\(C_l (1 + \alpha \sigma^2_l,\text{beam})\)”, where \(\sigma^2_l,\text{beam}\) is the noisiest beam eigenmode, as shown in Figure 12 of Hill et al. (2009). The 1D marginalized distributions for \(n_s\) are shown in the right panel of Figure 11. As with SRCMARG, the BEAMMARG result is indistinguishable from our standard result. Inflating the beam error by a factor of 20 results in a 14% increase in \(\delta n_s\), along with a slight shift in \(n_s\) away from unity.

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