K2 discovery of a circumsecondary disk transiting EPIC 220208795

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ABSTRACT

Context. Observations of the star EPIC 220208795 (2MASS J01105556+0018507) reveal a single, deep and asymmetric eclipse, which we hypothesise is due to an eclipsing companion surrounded by a tilted and inclined opaque disk, similar to those seen around V928 Tau and EPIC 204376071.

Aims. We aim to derive physical parameters of the disk and orbital parameters for the companion around the primary star.

Methods. The modelling is carried out using a modified version of the python package pyP1unS, and optimisation is done using Jupyter. The period analysis makes use of photometry from ground-based surveys, where we performed a period folding search for other possible eclipses by the disk. Parameters obtained by the best model fits are used to obtain the parameter space of the orbital parameters, while the most likely period obtained is used to constrain these parameters.

Results. The best model has an opaque disk with a radius of $1.14 \pm 0.03 R_\odot$, an impact parameter of $0.61 \pm 0.02 R_\odot$, an inclination of $77.01^\circ \pm 0.03^\circ$, a tilt of $36.81^\circ \pm 0.05^\circ$, and a transverse velocity of $77.45 \pm 0.05 \, \text{km s}^{-1}$. The two most likely periods are ~290 days and ~236 days, corresponding to an eccentricity of ~0.1, allowing us to make predictions for the epochs of the next eclipses. All models with tilted and inclined disks result in a minimum derived eccentricity of 0.3, which in combination with the two other known small transiting disk candidates V928 Tau and EPIC 204376071, suggest that there may be a common origin for their eccentric orbits.

Key words. eclipses – planets and satellites: rings – binaries: eclipsing

1. Introduction

Advances in high precision photometry has allowed astronomers to continuously monitor the apparent brightness of stars, and some of these stars exhibit deep and irregular eclipses in their apparent brightness over time. These patterns can come from intrinsic stellar variability (Joy 1945; Lanza et al. 2007; Olah et al. 2009; Handler 2013) or external objects transiting the star, ranging from exoplanets, to exocomets (Rappaport et al. 2018; Zieba et al. 2019), to material in and above circumstellar disks (Ansdell et al. 2019; Kennedy et al. 2020) or infalling material onto white dwarfs (Vanderburg et al. 2015; Gänsicke et al. 2019). Other light curves challenge an easy identification of the associated astrophysical processes - two notable cases are Boyajian’s star (Boyajian et al. 2016) and HD 139139 (Rappaport et al. 2019a). Dust that is believed to come from material in the inner part of circumstellar disks can produce dips in the brightness of stars that are called ‘dippers’ of ~10–50% (Ansdell et al. 2016, 2019; Alencar et al. 2010; Cody et al. 2014; Cody & Hillenbrand 2018). Circumstellar disks are a universal feature of star formation and also dictate the structure of the planetary system that can be formed by the material inside these disks (Williams & Cieza 2011). The formation of gas planets is thought to be through the accretion of material from circumstellar disks that subsequently passes through a circumplanetary disk. A tilted and inclined disk around a planet or substellar companion can cause a dip in stellar brightness inconsistent with a transiting exoplanet eclipse. The projection of such a disk creates an elliptical occulting region and (for non-zero impact parameters) creates an asymmetric eclipse.

Analysing such systems can yield insights into the formation mechanisms of planets, especially if the observed system is young, since planets form and grow within protoplanetary disks. The properties and dynamics of these disks are connected, not only to stellar evolution and the evolution of a circumstellar disk, but also to the properties and evolution of the planet (Armitage 2011; Kley & Nelson 2012). Analysing a transit can reveal the structure of the circumplanetary disk and provide insight into the dynamics and mechanics of ring and moon formation (Teachey et al. 2017). Examples of disk systems or dusty occultations analysed in previous studies are ‘J1407’ (V1400 Cen, Kenworthy & Mamajek 2015), EPIC 204376071 (EPIC 2043, Rappaport et al. 2019b), and V928 Tau (van Dam et al. 2020). In the case of J1407, a ring system is hypothesised to be around a secondary companion occulting the star, whereas the other two systems are of an inclined circular disk around a secondary companion occulting the star. An elliptical occulter has recently been suggested for the light curve seen towards the late type giant star labelled VVV-WIT-08 (Smith et al. 2021) discovered as part of the VISTA Variables in the Via-Lactea (VVV; Minniti et al. 2010; Minniti et al. 2017) survey, although this is closer in diameter to the J1407 occulter, and it produced an eclipse that lasted about 200 days.

The Kepler mission was launched to determine the frequency of Earth-sized planets in and near the habitable zone of Sun-like stars (Borucki et al. 2010). The mission has...
measured a large number of high-resolution light curves and has discovered a great deal of exoplanets (Thompson et al. 2018). After the failure of Kepler’s second reaction wheel in May 2013, the repurposed K2 mission (Howell et al. 2014) was able to survey the ecliptic plane, leading to numerous additional discoveries (Mayo et al. 2018). Several single, deep transit events have been identified in the K2 data (LaCourse & Jacobs 2018), and we examine the curves that have been classified as ‘Deep’, looking specifically for asymmetric eclipses as these hint at an elliptical occulter. EPIC 220208795 (EPIC 2202), also known by the aliases TIC 336889445, 2MASS J01105556+0018507, SDSS J011055.57+001850.6, and Gaia DR2 2534801707104852864, shows a deep and asymmetric transit event similar to the ones observed in Rappaport et al. (2019b); van Dam et al. (2020) and we model this eclipse with a tilted and inclined disk around a companion occulting the star. Section 2 gives a description of the telescopes and the photometry obtained from them, and the preliminary analysis on the K2 light curves to identify the most likely candidates for a tilted and inclined disk transit and why we proceeded with EPIC 2202. Section 3 describes the modelling of the asymmetric dip found in the EPIC 2202 light curve and we give a description of the best fits for two separate models. In Sect. 3.5 we perform a period folding analysis using ground-based survey photometry to determine the most likely period given the photometry available and identify other possible eclipses. Section 3.6 contains the orbital analysis of the best fit models, obtaining orbital parameters like the eccentricity and periastron distance. In Sect. 4 we discuss our results, point out limitations of our modelling, and compare them to two other disk systems from Rappaport et al. (2019b) and van Dam et al. (2020). Section 5 summarises our findings and presents suggestions for future research.

2. Data

The entire code used for the analysis and creating the figures in this paper is available online.

2.1. K2

The Kepler spacecraft is a 0.95 m Schmidt telescope with a 1.4 m diameter primary and a 110 deg field of view. The prime focus camera has 42 CCDs that are 2200×1024 pixels (Borucki et al. 2010). The extended K2 mission observed EPIC 2202 during campaign 8 (2016-Jan-03 to 2016-Mar-23) for a total of 79 days. It collected 3840 observations, with a cadence of 30 minutes, in the K_p filter, with a bandpass of 420–900 nm. Further information is listed in Table 1.

The light curve was extracted using EVEREST 2.0. (Luger et al. 2016, 2018), which is an open-source pipeline for removing instrumental noise from K2 data. It uses a variant of Pixel Level Decorrelation (PLD; Deming et al. 2015) to reduce systematic errors caused by the Kepler spacecraft’s pointing error, such as the 6 h trend in the raw aperture photometry, which compromises its ability to detect small transits. The light curve reveals a star with low amplitude flux variations (less than 1%), which were deemed small enough to avoid stellar variability modelling.

Although the main science mission for Kepler and K2 was the detection and characterisation of exoplanet transits, the library of light curves collected include many other types of astrophysical phenomena. To this end, LaCourse & Jacobs (2018) manually inspected 238399 stellar light curves and identified 48 stars that showed deep or unusual eclipses. Ten stars were observed on two separate campaigns resulting in a total of 53 light curves. The light curves were examined for asymmetry as this is a tell-tale sign of an elliptical occulter, which hints at circumsecondary disk transits. Of these targets we determined that EPIC 2202 was the most likely candidate for a tilted and inclined disk system, similar to the transits seen in the case of EPIC 2043 and V928 Tau. The eclipse of EPIC 2202 lasts for about 7.2 h with a maximum depth of about 25% (see Fig. 1).

2.2. TESS

The Transiting Exoplanet Survey Satellite (TESS) is a satellite designed to survey for transiting exoplanets among the brightest (and nearest) stars over most of the sky (Ricker et al. 2015). The TESS satellite orbits the Earth every 13.7 days on a highly elliptical orbit, scanning a sector of the sky spanning 24×96 deg for a total of two orbits, before moving on to the next sector. It captures images at a 2 s (used for guiding), 20 s (for 1 000 bright astroseismology targets), 120 s (for 200 000 stars that are likely planet hosts) and 30 min (full frame images) cadences. The instrument consists of 4 CCDs each with a field of view of 24×24 deg², with a wide band-pass filter from 600–1000 nm (similar to the I_c band) and has a limiting magnitude of about 14–15 mag (I_c). Further information is presented in Table 1.

EPIC 2202 was observed in Sector 3 (2018-Sep.-20 to 2018-Oct.-18) providing 770 photometric points with a cadence of 30 minutes and in Sector 30 (2020-Sep.-22 to 2020-Oct.-21) providing 3511 photometric points with a cadence of 10 min (see Fig. 2). Photometry for TESS has been processed using the eleonor (Feinstein et al. 2019), an open-source tool that produces light curves from TESS Full Frame Images.

2.3. Ground-based surveys

To supplement the space-based data obtained by K2 and TESS, several archival databases for ground-based surveys were queried for data on EPIC 2202, resulting in three sets of light curves. This relatively low number is most likely due to the fact that EPIC 2202 is faint (see Table 2).

The first of these ground-based surveys is the All Sky Automated Survey (ASAS; Pojmanski 1997, 2005; Simon et al. 2018). This is a survey consisting of two observing stations - one in Las Campanas, Chile, and the other on Maui, Hawaii. Recently each observatory was equipped with two CCD cameras using V and I filters and commercial f = 200 mm, D = 100 mm lenses, although earlier both larger (D = 250 mm) and smaller (50–72 mm) lenses were used. Most data are taken with pixel scale of ≈15″. ASAS splits the sky into 709 partially overlapping (9 × 9 deg² fields, taking on average 150 3-min exposures per night, leading to variable cadence of 0.3–2 frames per night. Depending on the equipment used and the mode of operation, the ASAS limiting magnitude varied between 13.5 and 15.5 mag in V, and the saturation limit was 5.5–7.5 mag. Precision is around 0.01–0.02 mag for bright stars and below 0.3 mag for the faint ones. ASAS photometry is calibrated against the Tycho catalogue, and its accuracy is not better than 0.05 mag for bright, non-blended stars.

The second is the All Sky Automated Survey for Supernovae (ASAS-SN; Shappee et al. 2014; Kochanek et al. 2017) which consists of six stations around the globe, with each station hosting four telescopes with a shared mount. The telescopes consist of a 14-cm aperture telephoto lens with a field of view of approximately 4.5×4.5 deg² and an 8.0” pixel scale. Two of
Table 1. Overview of the time-series photometric instruments.

| Survey     | Filter | $n_{tel}$ | $n_{phot}$ | Baseline (days) | Start date   | End date     | Field of view ($\text{deg}^2 \text{ cam}^{-1}$) | Pixel-scale ($\text{''} \text{ pix}^{-1}$) |
|------------|--------|-----------|------------|-----------------|--------------|--------------|-----------------------------------------------|--------------------------------------------|
| ASAS $^{(a)}$ | $I$    | 1–4       | 85         | 2026            | 27-12-2002   | 14-07-2008  | 4.8, 77.4                                    | 14.2                                       |
|            | $V$    | 829       | 7007       | 20-11-2000      | 27-01-2020   |              |                                               |                                            |
| ASAS-SN    | $g$    | 12        | 1604       | 1244            | 17-09-2017   | 12-02-2021  | 4.5                                          | 8.0                                        |
|            | $V$    | 8         | 1049       | 2371            | 02-06-2012   | 29-11-2018  |                                               |                                            |
| ATLAS      | $c$    | 2         | 356        | 2006            | 12-08-2015   | 07-02-2021  | 29.2                                         | 1.9                                        |
|            | $o$    |           | 892        | 1926            | 23-10-2015   | 01-02-2021  |                                               |                                            |
| K2         | $K_p$  | 1         | 3840       | 79              | 04-01-2016   | 23-03-2016  | 110                                          | 4.0                                        |
| TESS $^{(b)}$ | $I_C$  | 1         | 770        | 20              | 24-09-2018   | 14-10-2018  | 576                                          | 21                                         |

Notes. $^{(a)}$In 2000 the survey upgraded to two telescopes in Chile. In 2006 two telescopes in Hawaii were added. $^{(b)}$Centred on the traditional $I_C$ band, but has a 600–1000 nm bandpass.

The original stations (one in Hawaii and one in Chile) are fitted with $V$ band filters, whereas the additional stations (Chile, Texas, South Africa, and China) are fitted with $g$ band filters. ASAS-SN observes the whole sky every night with a limiting magnitude of about 17 mag in the $V$ and $g$ bands. The data was obtained through the publicly available Sky Patrol Service$^2$.

The third and final survey is the Asteroid Terrestrial-impact Last Alert System (ATLAS; Heinze et al. 2018; Tonry et al. 2018). This is a deep survey with a limiting magnitude of 19 and a precision of around 0.01–0.04 mag. It consists of two telescopes in Hawaii with a 0.65 m primary mirror and a 0.5 m Schmidt corrector. Mounted on the telescope is a camera with a 13.5-cm aperture lens, which when combined with the main telescope, provides a field of view of approximately 29 $\text{deg}^2$. The ATLAS survey uses three special filters $c$ (cyan, 420–650 nm), $o$ (orange, 560–820 nm) and $t$ (tomato, 560–975 nm), that are designed to be differentially sensitive to the silicate colours of stony asteroids. Data for the ATLAS survey was obtained via their online ‘forced photometry’ server$^3$.

Further information on the ground-based surveys is presented in Table 1 and a light curve of all the photometric time-series data collected is presented in Fig. 3.

2.4. Photometry and astrometric data
Table 2 summarises the available astrometry and photometry for EPIC 2202. The parallax and proper motion are listed from Gaia Early Data Release 3 (Gaia EDR3; Gaia Collaboration 2021), and the distance is a geometric estimate based on the Gaia EDR3 parallax taking into account a prior Galactic model (Bailer-Jones et al. 2021). Mean photometric magnitudes are provided by Gaia EDR3, the Two Micron All Sky Survey (2MASS; Skrutskie et al. 2006), and the Sloan Digital Sky Survey Data Release 8 (SDSS DR8; Aihara et al. 2011).

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$^2$ https://asas-sn.osu.edu

$^3$ https://fallingstar-data.com/forcedphot/
Fig. 2. Eleanor light curve of EPIC 2202. The upper panel shows Sector 3 and the lower panel shows Sector 30. We note that in Sector 30 there was significantly more scattered light, which explains the larger scatter and the larger absence of data between the two orbits.

Fig. 3. Time-series photometry of EPIC 2202 from several ground-based surveys as well as the K2 and TESS data. Photometry with a flux value greater than 150% or less than 10% was excluded, along with flux errors error greater than 100%.

3. Analysis
3.1. Stellar parameters
To independently estimate basic stellar parameters for EPIC 2202, we use the Virtual Observatory SED Analyzer 6.0 (VOSA; Bayo et al. 2008)\(^4\) to construct the star’s Spectral Energy Distribution (SED) and fit synthetic stellar spectra. For reddening, we adopt the full Galactic dust column in the direction of the star from Schlafly & Finkbeiner (2011) \((E(B-V) = 0.0219 \pm 0.0003)\) as the STILISM 3D reddening maps from Lallement et al. (2018) show that the dust in the direction of the star is mostly confined to within \(d < 335\) pc and the star lies at a Gaia EDR3-inferred distance of \(d = 487 \pm 6\) pc (Bailer-Jones et al. 2021). For dwarf stars of \(T_{\text{eff}} \approx 5000\) K, a typical Galactic ratio of total to selection extinction value of \(R_V = A_V/E(B-V)\) is 3.2 (McCall 2004), with \(\sigma_{R_V} \approx 0.18\) so we adopt an interstellar

\(^4\) http://svo2.cab.inta-csic.es/theory/vosa/
extinction $A_V = 0.070$. Using the relation $A_K = 0.382 E(B-V)$ from Bilir et al. (2008), we estimate the $K_s$ extinction to be $A_{K_s} = 0.0084$. Photometry from several surveys was included in the SED fitting, including: GALEX NUV (Bianchi et al. 2011), Sloan DR9 $ugriz$ (Adelman-McCarthy et al. 2011).

For the fits of the photometry to the synthetic stellar spectra (see Fig. 4), we constrained the extinction $A_V \in [0.06, 0.08]$, surface gravity $\log g \in [4, 5]$, and metallicity $[\text{M/H}] \in [-0.5, 0.5]$, as the star is clearly a dwarf and the photometric metallicities are approximately solar with $[\text{M/H}] = +0.02^{+0.15}_{-0.14}$ (Anders et al. 2019). For a grid of Kurucz ATLAS9 models (Castelli & Kurucz 2003), the best fit spectral template (fitting 27 photometric bands) was $T_{\text{eff}} = 5000 \, \text{K}$, $[\text{M/H}] = +0.5$, $\log g = 4$, $L = 0.407 \pm 0.010 \, L_\odot$. Varying whether the UV photometry (GALEX NUV) or redder IR bands (WISE W3 and W4) were included or not, and examining the scatter in output parameters when varying metallicity, surface gravity, or extinction within the previous bounds, all had negligible effect on the resultant bolometric flux and luminosity. No ultraviolet (NUV) or infrared (for example, WISE W3 and W4) excess was apparent. From this SED analysis we adopt the following parameters: bolometric flux $f_{\text{bol}} = 5.288(\pm0.130) \times 10^{-11} \, \text{erg} \, \text{s}^{-1} \, \text{cm}^{-2}$, apparent bolometric magnitude $m_{\text{bol}} = 14.194 \pm 0.027$ (IAU 2015 scale), absolute bolometric magnitude $M_{\text{bol}} = 5.716 \pm 0.036$ (IAU 2015 scale), bolometric luminosity $L = 0.4070 \pm 0.0137 \, L_\odot$, and $\log(L/L_\odot) = -0.3904 \pm 0.0146$ dex.

There are several other published $T_{\text{eff}}$ estimates for EPIC 2202, listed in Table 3. The $T_{\text{eff}}$ estimates are in reasonable statistical agreement with one another. The Anders et al. (2019) value is somewhat high, likely due to adopting an excessive extinction ($A_G = 0.20$) which does not appear to be supported by the extinction maps (Schlafly & Finkbeiner 2011; Lallement et al. 2018). Omitting the Anders et al. (2019) value, we adopt the unweighted mean $T_{\text{eff}}$ of the remaining values, and adopt the standard deviation as a conservative estimate of the uncertainty: $T_{\text{eff}} = 5060 \pm 50 \, \text{K}$ (1.0% uncertainty). This $T_{\text{eff}}$ is typical for a solar composition K2V star (Pecaut & Mamajek 2013)\(^5\). Combining the adopted $T_{\text{eff}}$ estimate with the previously estimated luminosity from our SED analysis, we estimate the radius of EPIC 2202 to be $R = 0.830 \pm 0.022 \, R_\odot$ (IAU nominal solar radius) or $577600 \pm 15100 \, \text{km}$ (2.6% uncertainty).

Only two mass estimates have been published since the availability of parallax data for EPIC 2202 by Gaia DR2. Anders et al. (2019) (StarHorse) estimates the mass to be $0.827^{+0.057}_{\text{−0.031}} \, M_\odot$, and the TIC version 9 (Stassun et al. 2019) estimates the

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**Table 2. Properties of EPIC 2202.**

| Property | Value | Ref. |
|----------|-------|------|
| $\alpha_{\text{CRS}}, \beta_{\text{CRS}}$ [hh mm ss] | 01:10:55.573 | 1 |
| $\delta_{\text{CRS}}$ [dd mm ss] | +00:18:50.52 | 1 |
| $\mu_\alpha$ [mas yr$^{-1}$] | 14.642 ± 0.026 | 1 |
| $\mu_\delta$ [mas yr$^{-1}$] | $-8.103 \pm 0.019$ | 1 |
| $\sigma$ [mas] | 2.0152 ± 0.0230 | 1 |
| Distance [pc] | 486.6$^{+4.7}_{-5.9}$ | 2 |
| $G$ [mag] | 14.2645 ± 0.0028 | 1 |
| $G_{\text{AP}}$ [mag] | 14.7271 ± 0.0033 | 1 |
| $G_{\text{SP}}$ [mag] | 13.6425 ± 0.0039 | 1 |
| $J$ [mag] | 12.897 ± 0.026 | 3 |
| $H$ [mag] | 12.431 ± 0.024 | 3 |
| $K_s$ [mag] | 12.321 ± 0.024 | 3 |
| $u$ (AB) [mag] | 17.63 ± 0.02 | 4 |
| $g$ (AB) [mag] | 16.096 ± 0.008 | 4 |
| $r$ (AB) [mag] | 16.69 ± 0.02 | 4 |
| $i$ (AB) [mag] | 14.123 ± 0.005 | 4 |
| $z$ (AB) [mag] | 14.832 ± 0.011 | 4 |
| $R$, $[R_\odot]$ | 0.830 ± 0.022 | 5 |
| $M$, $[M_\odot]$ | 0.85 ± 0.02 | 5 |
| $[\text{Fe/H}]$ [dex] | $+0.02^{+0.15}_{-0.14}$ | 6 |
| $\log g$ [$\log_{10}$ cm s$^{-2}$] | $4.56^{+0.05}_{-0.04}$ | 6 |
| $T_{\text{eff}}$ [K] | 5060 ± 50 | 5 |
| $f_{\text{bol}}$ [$10^{-11}$ erg s$^{-1}$] | 5.288 ± 0.130 | 5 |
| $m_{\text{bol}}$ [mag] | 14.194 ± 0.027 | 5 |
| $M_{\text{bol}}$ [mag] | 5.716 ± 0.036 | 5 |
| $L_{\text{bol}}$ [$L_\odot$] | 0.4070 ± 0.0137 | 5 |
| $\log(L/L_{\text{bol}})$ [dex] | $-0.3904 \pm 0.0146$ | 5 |

**Table 3. Estimated effective temperatures for EPIC 2202.**

| $T_{\text{eff}}$ [K] | Reference |
|---------------------|-----------|
| 4997$^{+97}_{-19}$ | Gaia Collaboration (2018) |
| 5000 | This work |
| 5064$^{+102}_{-101}$ | Huber et al. (2016) |
| 5090 ± 49 | Bai et al. (2019) |
| 5095 ± 122 | Stassun et al. (2019) |
| 5105 ± 138 | Hardegree-Ullman et al. (2020) |
| 5206$^{+156}_{-104}$ | Anders et al. (2019) |
| 5060 ± 50 | Adopted |

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\(^5\) [http://www.pas.rochester.edu/~emamajek/EEV_dwarf_UBVIColors_Teff.txt]
mass to be 0.850 \, M_\odot \, (\text{no uncertainties}). Huber et al. (2016) estimated the mass to be 0.816^{+0.050}_{-0.070} \, M_\odot, however this was without the benefit of a trigonometric parallax. Using the most recent mass-luminosity calibration for binary stars with dynamical masses from Eker et al. (2018), our luminosity estimate \( \log (L/L_\odot) = -0.390 \) translates to a mass estimate of 0.858 \, M_\odot. Based on these three estimates (Stassun et al. 2019; Anders et al. 2019, and this work), we adopt a stellar mass of 0.85 \pm 0.02 \, M_\odot.

### 3.2. Eclipse modelling

The eclipse observed in K2 photometry is most likely due to an occulter transiting EPIC 2202. The asymmetry and depth of the eclipse, as well as the short duration, make it unlikely for the transiting object to be a spherically symmetric object. We assume that the eclipse is due to a tilted and inclined disk gravitationally bound to an unseen secondary companion that creates an elliptical occulter due to its projected geometry. We model the light curve with two models – a hard edged disk, and a two ring disk consisting of an inner opaque disk and an outer, partially transmissive ring.

#### 3.2.1. Model parameters

The modelling is performed using a modified version of the pyPp1usS package that models the transit of an oblate exoplanet or an exoplanet with rings by calculating the area hidden by the occulter across the limb-darkened disk of the star (Rein & Ofir 2019). We use this to model the light curve based on the position and geometry of the disk with respect to the (limb-darkened) host star. The code works the spatial domain (in units of stellar radii) and geometry of the disk with respect to the (limb-darkened) host star. The code calculates the spatial domain (in units of stellar radii) and geometry of the disk with respect to the (limb-darkened) host star.

To calculate a lower bound on \( v_t \), we use the method described in van Werkhoven et al. (2014). This method requires the linear limb-darkening parameter, \( u \), of the star, which has been determined using the jktebop programme written by Southworth (2015). This programme uses the surface gravity, \( \log g \), the effective temperature, \( T_{\text{eff}} \), and the metallicity, \( \text{[Fe/H]} \), of the star to linearly interpolate the tables from Sing (2010) to calculate \( u \) for the Kepler bandpass. These values are derived by Huber et al. (2016) for 138 600 K2 targets in campaigns 1–8, using proper motions, colours, parallaxes, spectroscopy, and stellar population models. For EPIC 2202 we have \( \log g = 4.591, T_{\text{eff}} = 5064 \, \text{K}, \) and \( \text{[Fe/H]} = -0.112, \) which results in a value of \( u = 0.6681 \).

Using \( u \), the steepest time gradient of the light curve, \( L \), and taking \( R = R_\star \), we can get a lower limit on \( v_t \) following the method of van Werkhoven et al. (2014):

\[
v_t = L R_\star \left( \frac{2u - 6}{12 - 12u + 3\pi u} \right).
\]

For our obtained value of \( u \) and for \( L = 3 \, L_\star \, \text{day}^{-1} \) we obtain a lower limit of \( v_t = 4.3 \, R_\star \, \text{day}^{-1} \).

The free parameters for the model are the radius of the disk, \( R_d \), the impact parameter (perpendicular distance between the centre of the star and the orbital path), \( b \), the inclination of the disk, \( i \), the tilt of the disk (the angle with respect to the orbital path), \( \phi \), the transverse velocity of the disk, \( v_t \), the time shift with respect to the time of closest approach, \( \delta t \), and the opacity of the disk \( \tau \).

We use emcee (Foreman-Mackey et al. 2013) to explore these parameters and use Markov chain Monte Carlo (MCMC) methods to determine the best fit of the models. To properly explore this parameter space physical bounds on these parameters must be applied. These are as follows: an upper limit for \( R_d \) of ten times the radius of the star, \( R_\star \), the upper and lower limit on \( b \) are such that the disk must transit the star, the bounds on \( i \) are from 0° (face-on disk) to 90° (edge-on disk), \( \phi \) has been limited from 0° to 90° because of reflection symmetries induced by the combination of \( b \) and \( \phi \), and a range of \( \delta t \) between −10 and 10 days. We take \( v_t \), between 4.3 \, R_\star \, \text{day}^{-1} \) and 20 \, R_\star \, \text{day}^{-1}, where the upper bound has been deemed large enough and the lower bound has been calculated above. The opacity of the disk, \( \tau \), can generally be bound between 0 and 1, but here we fix \( \tau \) at 1, to get the smallest possible disk.

#### 3.2.2. Hard edged disk

The first model we explore is a hard edged disk. We fix the opacity such that the disk is completely opaque, since this gives the smallest possible disk diameter for a given light curve gradient. The smallest physically plausible disk model then gives a lower bound on the mass of the secondary companion through the size of the companion’s Hill sphere. We are interested in the smallest possible disks as these are deemed more likely to exist due to stability of the disk, and orbital velocity considerations.

To determine a suitable starting point of the fit we set a Gaussian prior bound by the limits in Table 4 and described in the previous section. We ran this for 600 links with 1000 walkers. We used a local minimum from this parameter investigation to set the initial starting point for the walkers of the extended optimisation. This starting point was expanded with Gaussian priors to produce the initial parameter spread summarised in Table 5.

| Parameter | Lower bound | Upper bound |
|-----------|-------------|-------------|
| \( R_d \) [\( R_\star \)] | 0.9 | 1.4 |
| \( τ \) [-] | −0.04 | −0.04 |
| \( τ_e \) [-] | − | − |

| Parameter | Median | Spread |
|-----------|--------|--------|
| \( R_d \) [\( R_\star \)] | 0.9 | 0.1 |
| \( τ \) [-] | −0.04 | 0.01 |
| \( τ_e \) [-] | − | − |

| Parameter | Hard edged disk | Soft edged disk |
|-----------|-----------------|-----------------|
|\( R_d \) [\( R_\star \)] | 0.9 | 0.1 |
| \( τ \) [-] | −0.04 | −0.04 |
| \( τ_e \) [-] | − | − |
Based on the one percent residuals that we see in the hard edged model, we added an outer ring with variable thickness, \( t \).

The residuals of the fit are at around the 1 percent level. However, it is clear from the systematic deviations of the residuals that the modelled eclipse is not wide enough at the start and end of the eclipse and not deep enough at maximum occultation to follow the measured light curve. This pattern suggests that the disk is not quite large enough and prompts an extension of the hard edged model.

### 3.3. Soft edged disk

Based on the one percent residuals that we see in the hard edged model, we added an outer ring with variable thickness, \( t \), and opacity, \( \tau_e \), and centre the prior of the soft edged disk model on the best fit of the hard edged disk model. We set the bounds of \( \tau_e \) to be between 0 and 1, and those of \( t \) to be the same as \( R_d \).

The residuals of the fit are at around the 1 percent level. However, it is clear from the systematic deviations of the residuals that the modelled eclipse is not wide enough at the start and end of the eclipse and not deep enough at maximum occultation to follow the measured light curve. This pattern suggests that the disk is not quite large enough and prompts an extension of the hard edged model.

#### 3.4. Comparison of the two models

The residuals from both models are similar in amplitude and shape, as can be seen in Fig. 7, which overlays the best fits. The soft edged model does give a slightly lower \( \chi^2 \) value (65 955 for 72 photometric points) compared to the hard edged model (70 729 for 72 photometric points), which implies that the soft edged model is a slightly better fit to the data. However, it is also clear that the addition of a soft edge to the hard edge model was not enough to reduce the observed residuals and thus does not explain their origin.
3.5. Period folding

To search for a plausible orbital period of an occulter orbiting EPIC 2202, we supplemented the data from K2 with data from TESS and several ground-based surveys, as can be seen in Fig. 3.

We folded all the data over a large (60–1800 days) high temporal resolution (step size of 0.001 days) period grid and examined the folded light curve in the region where the K2 eclipse occurs. The routine flags periods where at least three photometric points lie within three times the photometric error of that point from the linearly interpolated K2 data.

The best period candidates were determined by comparing the \( \chi^2 \) value of folded data with respect to the K2 data (\( \chi^2_{\text{model}} \)) to the \( \chi^2 \) value of the folded data with respect to a flat line (\( \chi^2_{\text{flat}} \), i.e. a no eclipce model). The lower the ratio \( \chi^2 - \text{ratio} = \chi^2_{\text{flat}}/\chi^2_{\text{model}} \), the more the folded data follows a no-eclipse scenario. The higher the \( \chi^2 - \text{ratio} \), the more likely the folded data follows an eclipse scenario. A further criterion was that at least one point of the folded data lies in the phase interval where the flux of the K2 data is below a flux of 85%. This is to ensure that the periods investigated actually contain data in the deepest part of the eclipse.

Of all the periods investigated, 13 out of 15 periods with a \( \chi^2 - \text{ratio} > 400 \) suggest a fundamental period of 290.230 days, and 2 out of 15 periods suggest a fundamental period of 235.587 days (see Table 7 for more details). For a period of 290.230 days, photometry with a \( \chi^2 - \text{ratio} > 400 \) was found for 2 and 4 times the period, but not for 3 times the period due to incomplete photometric coverage (see Fig. 8). For a period 235.587 days there was no reliable photometry for a multiple of the period, due to incomplete coverage or due to the absence of points in the deepest part of the eclipse (with flux values below 85%).

Note that any of the periods where no photometry was folded into the eclipse cannot be ruled out by this period folding analysis. Therefore, there are many more possible periods for EPIC 2202 that cannot be investigated with the available data.

Considering the most likely shorter orbital periods (see Table 7), we can make predictions for the upcoming eclipses of EPIC 2202. The dates, times, and relation to the observed eclipse for the upcoming four eclipses for each period are listed in Table 8.

| Period (day) | \( \eta_{\text{phot}} \) (within eclipse) | \( \chi^2 - \text{ratio} \) | \( \chi^2_{\text{flat}}/\chi^2_{\text{model}} \) |
|-------------|---------------------------------|----------------|-----------------|
| 290.230     | 4                               | 423.564        |                  |
| 235.587     | 3                               | 410.788        |                  |

Notes. Grey rows correspond to \( P_{\text{orb}} = 235.587 \) days. white rows correspond to \( P_{\text{orb}} = 290.230 \) days.

3.6. Orbital parameters

We have obtained two transverse velocities from our hard edged and soft edged model, 11.6 and 11.5 \( R_\odot \) day\(^{-1} \) respectively, and taking \( R_\star = 0.83 R_\odot \), this corresponds to 77.4 and 76.9 km s\(^{-1} \). The eclipse observed by K2 does not repeat within the observation window of K2, which gives us a lower limit on the orbital period, \( P_{\text{orb}} \). \( P_{\text{orb}} \) must be larger than the longest time within the K2 data where no eclipse appears, which in this case is 60 days.

If we take the modelled transverse velocity, \( v_t \), and calculate the period of a circular orbit, it results in \( P_{\text{orb}} = 25 \) days, which is below the minimum period limit set by the rest of the K2 data, meaning that a circular orbit with this velocity is ruled out. We thus explore eccentric orbits, where we assume that \( v_t = v_{\text{peri}} \), the velocity at periastron, and create a grid consisting \( P_{\text{orb}} \), and the mass of the companion, \( M_p \). We follow the analysis as performed by van Dam et al. (2020) in Sect. 4.4, which we describe shortly here. Using Kepler’s Third Law we use \( P_{\text{orb}} \) to obtain the...
semi-major axis, $a$ – when combined with $e_{\text{peri}}$ and $M_p$, these then determine the eccentricity, $e$, using the vis-viva equation. Finally as a stability criterion, we state that $M_p$ must be large enough that the Hill radius, $R_{\text{Hill}}$, of the companion can support the circumsecondary disk ($R_{\text{disk}} < 3R_{\text{Hill}}$). We use the same $M_p$ bounds from 0–80 $M_J$, with the upper limit chosen as an inclusive H-burning limit (Saumon & Marley 2008), and a more conservative limit of 73–74 $M_J$ (Baraffe et al. 2015; Forbes & Loeb 2019). For $P_{\text{peri}}$, we set limits of the grid from 60–1800 days, set the upper limit is considered arbitrarily large and is related to the fact that the available photometry does not extend further away from the observed eclipse. The results for the hard edged and soft edged disk models are indistinguishable due to their similar $e_{\text{peri}}$. The parameter space investigation is depicted in Fig. 9 and the orbital parameter ranges at the most likely periods (from Table 7) are presented in Table 9.

The minimum mass is about 1.5 $M_J$, set by the Hill sphere stability criterion ($R_{\text{disk}} < 0.3R_{\text{Hill}}$) during the periastron passage of the disk. For the most likely period of 290.230 days, the eccentricity of the orbit is around $e = 0.72$, but as can be seen in Fig. 9, the minimum eccentricity for the system can be as low as $e = 0.36$. Longer orbital periods require higher eccentricities in order to keep the velocity at periastron equal to the derived transverse velocity.

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**Fig. 9.** Parameter space maps for a period between 60 and 1800 days, and mass between 0 and 80 $M_J$. The left side is the hard edged disk, the right side is the soft edged disk. The white space at the bottom of the plots is from the constraint $R_s < 0.3 r_1$. 

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Using the most likely periastron distance we found, and the effective temperature of the star, we can make an estimate of the equilibrium temperature of the disk. For a periastron distance of 0.23 AU, corresponding to a period of about 290.230 days, the equilibrium of a fully absorbing disk around a companion orbiting EPIC 2202 would be about 420 K. Using the second period of 235 days would result in a lower periastron distance, and thus a higher equilibrium temperature, of about 430 K. These temperatures exclude the possibility of a disk made of ice, so we conclude the disk is made of silicates.

### 4. Discussion

#### 4.1. The EPIC 2202 eclipser

We have found two likely models of disks that could be responsible for the eclipse of EPIC 2202, a fully opaque disk, and an opaque disk with a soft edge. Taking the simpler of the two models, the hard edged model gives a simple template that can be improved upon by adding more degrees of freedom, such as adding scattering or modelling an exponentially thinning disk. We have found a most likely period for the system based on time-series photometry from ground-based surveys, but it does not cover the whole time domain, which, combined with the short duration of the eclipse likely means that we are missing other eclipses that could determine the period of the system.

#### 4.2. A population of small disk occulters

There have been two other cases of circumsecondary disk occulters similar in size to EPIC 2202. The first is a disk around EPIC 2043 (Rappaport et al. 2019b), and the second is a disk around the binary star system V928 Tau (van Dam et al. 2020). A comparison between the radii of these stars as well as the parameters of the disk model obtained can be seen in Table 10, and a comparison of the three systems with disk models can be seen in Fig. 10. Note that for EPIC 2202, Table 10 and Fig. 10 show the simpler hard edged model.

**Table 9.** Orbital parameters for most likely periods.

| Parameter | $P_{\text{orb}} = 290$ days | $P_{\text{orb}} = 235$ days |
|-----------|-----------------------------|-----------------------------|
| $r_{\text{ap}}$ [AU] | 1.40–1.44 | 1.20–1.21 |
| $r_{\text{pen}}$ [AU] | 0.22–0.24 | 0.22–0.23 |
| $r_{\text{H}}$ [AU] | 0.02–0.07 | 0.02–0.07 |
| $e$ [-] | 0.72–0.73 | 0.67–0.69 |

**Table 10.** Best fit values for the small disk occulter systems.

| Parameter | EPIC 2202 (a) | EPIC 2043 | V928 Tau (b) |
|-----------|---------------|-----------|--------------|
| $R_*$ [R$_\odot$] | 0.83 | 0.63 | 1.38 |
| $R_\delta$ [R$_\odot$] | 1.37 | 4.20 | 0.99 |
| $R_\beta$ [R$_\odot$] | 1.14 | 2.60 | 1.39 |
| $b$ [R$_*$] | -0.74 | -0.88 | -0.25 |
| $i$ [°] | 77.0 | 77.8 | 56.8 |
| $\phi$ [°] | -36.8 | 18.1 | 41.2 |
| $v_0$ [km s$^{-1}$] | 77.4 | 38.0 | 73.5 |
| Min. $e$ [-] | 0.36 | 0.33 | 0.30 |
| Min. $r_{\text{pen}}$ [AU] | 0.17 | 0.13 | 0.10 |

**Notes.** (a)Values from hard edged model. Note that $b$ and $\phi$ are flipped for easier comparison between the systems. (b)Values taken around the primary star.
Future research includes constraining the period of the companion by detecting other eclipses, or by spectroscopic monitoring to determine the mass, or upper limit on the mass, of the secondary companion. The star is an early K dwarf star, which is more amenable to spectroscopic radial velocity measurements, as the other two systems are around M dwarf stars. If we assume that the periastron velocity was observed during the eclipse, we can estimate the amplitude of the radial velocity curve to be of the order of 70 km s$^{-1}$ ($M_p/M_2 \approx 250$ m s$^{-1}$), which would be possible to observe with a spectrograph on an 8m-class telescope. The eclipse duration is only 7.2 h, which makes it difficult to detect from ground-based observatories subject to diurnal window functions, but using the predictions for next eclipses based on the found periods, there could be observations planned for these dates to constrain the period of the companion. An observational campaign for the American Association of Variable Star Observers (AAVSO) would be a likely path for detecting subsequent eclipses. All three light curves are more rare for multiple exoplanet systems (Limbach & Turner 2015). This minimum eccentricity could be explained by a perturbing companion that causes the companion with the inclined disk to be in an eccentric orbit – the V928 Tau system is a binary star system, but we are not currently aware of any companions around EPIC 2043 or EPIC 2202. A perturbing companion may provide an explanation for the disks in the form of a planetesimal collision caused by a strong scattering event that put the companion into the observed eccentricity that we deduce.

All three light curves show a similar pattern of residuals – the points of ingress and egress show deviations from our model leading to the large $\chi^2$ values, and there is a difference in the deepest part of the eclipse, possibly as a result of the model not constraining enough degrees of freedom for an accurate model. Especially in the case of EPIC 2043 and V928 Tau, there is a distinct and similar down-up-down-up-down pattern in the residuals (see Figs. 5–7). For EPIC 2043, the residual pattern starts with a down-up pattern, but afterwards it is less distinct. The down-up-down-up-down pattern comes from the actual eclipse having less steep ‘wings’ than the model allows for and a deeper minimum. Both for EPIC 2202 and V928 Tau adding a soft edge to the model did not significantly improve the fit. This alludes to a physical phenomenon which is not included in the model. Future studies will include forward scattering and more complex structures for the putative disks.

5. Conclusions

We looked at the light curves initially found in K2 data by LaCourse & Jacobs (2018), from which we selected EPIC 2202 as the most likely candidate of a circumsecondary disk transit. Using a modified version of pyPluSS in combination with emcee, we modelled the eclipse with a hard edged and soft edged disk. Both of the models have similar residuals that the fitting did not account for. This distinct down-up-down-up-down pattern in the residuals is also seen in the proposed inclined disk system of V928 Tau, and also partly in EPIC 2043. The hard edged model is the simplest of the two, and so we keep our model of an opaque disk around a companion of EPIC 2202, with a disk size of 1.37 $R_*$, an impact parameter of 0.74 $R_*$, an inclination of $77.01^\circ$, a tilt of 36.81$^\circ$ and a transverse velocity of 77.4 km s$^{-1}$. A period analysis using ground-based data of EPIC 2202 resulted in two most likely periods, namely, $235.587$ and $290.230$ days. Taking the fitted transverse velocity we derived orbital parameters of the companion, such as a minimum eccentricity of $-0.36$. EPIC 2043 and V928 Tau have roughly the same minimal eccentricity. This eccentricity could be caused by a perturbing companion, but for EPIC 2043 and EPIC 2202 it is not known if there is such a perturbation.

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