The unusual X-ray light-curve of GRB 080307: the onset of the afterglow?

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ABSTRACT

Swift-detected GRB 080307 showed an unusual smooth rise in its X-ray light-curve around 100 seconds after the burst, at the start of which the emission briefly softened. This ‘hump’ has a longer duration than is normal for a flare at early times and does not demonstrate a typical flare profile. Using a two component power-law-to-exponential model, the rising emission can be modelled as the onset of the afterglow, something which is very rarely seen in Swift-X-ray light-curves. We cannot, however, rule out that the hump is a particularly slow early-time flare, or that it is caused by upscattered reverse shock electrons.

Key words: gamma-rays: bursts — X-rays: individual (GRB 080307) — radiation mechanisms: non-thermal

1 INTRODUCTION

Although many Gamma-Ray Burst (GRB) X-ray light-curves roughly follow a ‘canonical’ decay of initially steep and then flat (the so-called ‘plateau’ phase), before settling into a classical power-law decline phase (see, e.g. Nousek et al. 2006; Zhang et al. 2006), some are noticeably different. Examples include GRB 060105, which had an unusually shallow initial decay (Godet et al. 2008); GRB 060218, which was clearly associated with supernova 2006aj and showed the break-out of a shock wave from the exploding progenitor (Campana et al. 2006); GRB 061007, where the X-ray decay continued with no breaks until at least 10⁶ seconds after the trigger (Schady et al. 2007; Mundell et al. 2007); GRB 070110, which showed an abrupt drop in the X-ray emission at the end of the ‘plateau’ stage (Troja et al. 2008). It is examples such as these, whose behaviour differs from those of more typical afterglows, which provide further insight into the physics of the bursts, allowing the testing of models for the early stages of GRB afterglows.

GRB 080307 has a very unusual X-ray light-curve, reminiscent of GRB 060218. The emission rose smoothly until about 200 seconds after the trigger, at which point a long, steady decay set in, continuing unbroken until the afterglow could no longer be detected after a few 10⁵ s. Here we show that, in contrast to GRB 060218, the early emission peak is probably the rising afterglow emission; there is no evidence for a supernova shock breakout.

The observations are described in Section 2 with details on the Swift spectral analysis and results in Sections 3.1 and 3.3. Ground-based optical follow-up findings are given in Section 3.4. Section 4 presents a discussion.

2 OBSERVATIONS

2.1 Swift

Following the BAT trigger at 11:23:30 UT on 2008-03-07, Swift (Gehrels et al. 2004) immediately slewed to the burst
and started collecting X-ray Telescope (XRT; Burrows et al. 2005) and UV/Optical Telescope (UVOT; Roming et al. 2005) data at T+99 s and T+105 s respectively. The XRT centroided onboard, though a refined position was quickly determined from the promptly-downlinked SPER (Single Pixel Event Report) data (Holland et al. 2008). The most accurate Swift position was later calculated by using the XRT-UVOT alignment and matching UVOT field sources to the USNO-BI catalogue (see Goad et al. 2007b for more details): \textit{R.A. = 09\textdegree 06'30.72, Dec. = +35\textdegree 08'20.3'}. The data were processed with the standard Swift FTools, using v2.8 of the software (release date 2007-12-06) and the corresponding calibration files. The XRT stayed in Windowed Timing (WT) mode until around 614 s after the trigger, at which point the count-rate dropped below 2 count s\(^{-1}\) and the XRT automatically switched to Photon Counting (PC) mode. The PC data obtained towards the end of the first orbit were piled-up, so the light-curve and spectrum were extracted using an annulus to exclude the central four pixels of the Point Spread Function [PSF; the exclusion radius was estimated by fitting the wings of the PSF with the known King function parameters (Moretti et al. 2005) and determining at which point the observed profile deviated from the expected model]. A comparison of ancillary response files with and without the PSF correction can be seen in Table 1; these have not been corrected for the galactic nucleus. The afterglow itself was not detected during this period.

### 2.2 Ground-based follow-up

Optical observations were made with GMOS (Gemini Multi-Object Spectrograph) on Gemini North (Tanvir 2008), the Auxillary Port Camera on the William Herschel Telescope (WHT/AUX) and the 2-m Faulkes Telescope South (FTS). UKIRT/WFCAM (UK Infrared Telescope/Wide Field Camera) also observed and detected the afterglow in the IR; in this case, a sequence of exposures was obtained using an annulus to exclude the central four pixels of the Point Spread Function [PSF; the exclusion radius was estimated by fitting the wings of the PSF with the known King function parameters (Moretti et al. 2005) and determining at which point the observed profile deviated from the expected model]. A comparison of ancillary response files with and without the PSF correction can be seen in Table 1; these have not been corrected for the galactic nucleus. The afterglow itself was not detected during this period.

The optical and near-IR magnitudes measured are provided in Table 1 these have not been corrected for the reddening of \(E_{B-V} = 0.03\) mag (Schlegel, Finkbeiner & Davis 1998). UKIRT magnitudes were calibrated against 15 2MASS stars in the field. The WHT and FTS-\(i'\) data were calibrated using the SDSS system (Cool et al. 2008), while the USNOB1-0 catalogue was used for the \(B\) and \(R\) Bessel filters (Monet et al. 2003).

The Gemini data revealed a faint point source within the revised XRT error circle (Goad et al. 2008), providing an accurate position of \(\text{R.A. = 09\textdegree 06'30.30, Dec. = +35\textdegree 08'20.1'}\) (J2000; uncertainty of 0.2 arcsec) for the optical afterglow. Figure 1 shows the Gemini image (S1 is the position of the afterglow; S2 is the contaminating source discussed below).

### 2.3 Chandra

During an initial inspection of the XRT data, the X-ray light-curve showed a break at around \(4 \times 10^4\) s after the trigger. An apparent very flat decay then continued until the end of the Swift observations, around \(6 \times 10^6\) s (Holland et al. 2008b). Because Gemini optical imaging showed the presence of a nearby object (Figure 1), just a few arcseconds away from the GRB position, we obtained a late-time Chandra Target of Opportunity observation in September 2008, 6.5 months after the burst, to determine whether the slow X-ray decline was intrinsic or due to a contaminating source.

The Chandra data did reveal that there was an X-ray source coincident with the optical one, well within the point-spread function of the XRT. This contaminating source is at a position of \(\text{R.A. = 09\textdegree 06'31.04, Dec. = +35\textdegree 08'22.4'}\) (J2000), 3.8 arcsec from the afterglow position (determined from the optical data). Its \(i'\)-band magnitude is \(\sim 21.85\) \([J = 20.67, H = 19.88, K = 18.93]\) from UKIRT observations, with uncertainties of \(\sim 0.16\) mag, and the X-ray flux is approximately \(2.7 \times 10^{-14}\) erg cm\(^{-2}\) s\(^{-1}\) (observed; \(3.8 \times 10^{-14}\) erg cm\(^{-2}\) s\(^{-1}\) unabsorbed). This implies an optical-to-X-ray ratio of \(\alpha_{\text{ox}} \sim 1.4\), typical for an active galactic nucleus. The afterglow itself was not detected during the Chandra pointing.

The flux of the nearby source corresponds to \(\sim 8 \times 10^{-4}\) count s\(^{-1}\) in the XRT. This value has been subtracted off all the X-ray data points in the light-curve.

At the flux level of the contaminating source, the sky density of X-ray sources is around 57 per square degree (Mateos et al. 2008). This gives approximately a 1% chance of finding such a source in the extraction region used (radius of 28 arcsec). Given that Swift has detected more than 350 GRBs to date, it is likely that other X-ray light-curves reaching similar levels have been affected by contaminating sources. A possible example, noted by the authors, is...
GRB 050422 (Beardmore et al. 2007), which also showed a levelling-off of the light-curve at later times.

3 ANALYSIS AND RESULTS

3.1 Swift-BAT

The Burst Alert Telescope (BAT; Barthelmy et al. 2005) data give a $T_{90}$ value of 126 ± 26 s. A spectrum covering this time can be fitted with a single power-law photon index of $\Gamma = 1.81^{+0.24}_{-0.22}$, with no requirement for any form of energy cut-off or break. The corresponding 15–150 keV flux is $6.3 \times 10^{-9}$ erg cm$^{-2}$ s$^{-1}$.

Figure 2 shows the BAT light-curves over the standard energy bins (15–25, 25–50, 50–100 and 100–350 keV), while Figure 3 demonstrates how the $\gamma$-ray emission softened during the burst. Fitting spectra for $T = 0–50$ s and 50–126 s also shows this trend: $\Gamma_{0–50s} = 1.53 \pm 0.18$ and $\Gamma_{50–125s} = 3.18^{+0.96}_{-0.70}$.

Figure 2 also shows how the peak of the $\gamma$-ray data both moved to later times and broadened at softer energies; this trend continued for the X-ray data. A Gaussian component was applied to each light-curve to parametrize the changes in a simple manner and the results are given in Table 2.

3.2 X-ray analysis

3.2.1 Temporal Analysis

Excluding the hump in the light-curve and fitting the XRT data after 800 s, a single power-law is a good fit, with a decay slope of $\alpha = 1.95 \pm 0.10$ (where $F_{\nu,t} \propto \nu^{-\alpha} t^{-\Gamma}$ and the photon spectral index, $\Gamma = \beta + 1$). After about $2 \times 10^5$ s, the afterglow is no longer significantly detected above the nearby source. A slope of $\alpha \sim 1.95$ is more typical of the ‘normal’ decay phase of the ‘canonical’ X-ray afterglow (Nousek et al. 2006; Zhang et al. 2006; Evans et al. 2009; Racusin et al. 2009), being significantly steeper than the ‘plateau’ stage.

The fluence of the hump emission between 100–600 s is $\sim 2.1 \times 10^{-7}$ erg cm$^{-2}$, about a quarter of the time-averaged BAT (15–150 keV) burst fluence measured by Sato et al. (2008).

Figure 4 compares the hard (1.5–10 keV) and soft (0.3–1.5 keV) light-curves, showing there is a softening during the...
beginning of the brightening phase, after which the hardness ratio remains close to constant; the data beyond 600 s show no evidence for a change in hardness either. This early behaviour is unusual for a GRB ‘flare’ or ‘pulse’: typically, as the count-rate increases, so, too, does the hardness (e.g., Golenetskii et al. 1983; Ford et al. 1995; Borgonovo & Ryde 2001; Goad et al. 2007a; Page et al. 2007). However, the start of this rise may not have been caught by the XRT, so there could have been an earlier hardening – and some flare does seem to start hard and then just soften (e.g., GRB 050502B, Falcone et al. 2006). The spectral evolution of flares does tend to continue during the peak and beyond, though, which is not seen in this case; after about 140 s post-trigger, there is little or no further change of spectral shape in GRB 080307.

### 3.2.2 Spectral Analysis

Time-sliced spectra were extracted throughout the hump in the light-curve from \( T_0 + 100 - 600 \) s, and the resulting power-law fits are shown in the bottom panel of Figure 4. The absorbing column was fixed at the Galactic value \((2.37 \times 10^{20} \text{ cm}^{-2}); \) Kalberla et al. 2005, since no higher \( N_H \) was suggested by the fit. An alternative fit of keeping the power-law index fixed and allowing the variation of \( N_H \) to account for the difference in spectral shape did not provide such a good fit (\( \chi^2 \) was 10 higher, for one fewer degree of freedom).

A spectrum extracted for the PC data from \( T_0 + 4.4 \) ks onwards (to avoid any residual emission from the ‘flare’) has \( \Gamma = 1.66 \pm 0.30 \). Note, however, that the fit is improved if \( N_H \) is allowed to vary (99.5% via the F-test), leading to \( \Gamma = 2.16^{+0.36}_{-0.33} \), with a total column at \( z = 0 \) (including the Galactic value) of \( N_H = (1.8^{+2.3}_{-1.5}) \times 10^{21} \text{ cm}^{-2} \). This is discussed further in Section 4.

### 3.3 Swift-UVOT

No afterglow was detected by the UV/Optical Telescope (UVOT; Roming et al. 2005). Holland (2008) lists the 3\( \sigma \) upper limits, which range between \( > 20.6 - 21.4 \) for the optical and UV filters (summed over a few hundred to a few thousand seconds after the trigger). The afterglow was also not detected in the broad-bandpass white filter, to a limiting magnitude of \( > 22.3 \).

### 3.4 Ground-based follow-up

Gemini detected a faint \( i \approx 22.3 \) point source approximately 90 minutes post-trigger (Figure 1), which faded by about 0.1 mag during the course of this sequence (Table 1). WHIT also found a faint blue source approximately 11 hours post-burst in \( r \) and \( i \) filters, although subsequent observations with the WHIT over the course of several weeks showed no significant variation in the luminosity of the source, indicating the host galaxy was being detected.

Figure 5 plots the optical and IR data listed in Table 1. The second UKIRT \( H \)-band measurement seems to be slightly in excess of the previous point, and the \( J \)-band point around 4.2 ks is above that at \( \sim 3.8 \) ks. Unfortunately there are no simultaneous X-ray observations, but the X-ray data at the beginning of the second orbit (\( \sim 4.4 \) ks) do seem to be slightly higher than the underlying power-law. It is therefore plausible that there was a similar fluctuation in the X-ray and IR bands around this time.

### 4 DISCUSSION

#### 4.1 Limits on a supernova component

The late time limit on variability between the final two \( r \)-band observations is \( 0.23 \pm 0.22 \) \( \mu \)Jy, corresponding to a 3\( \sigma \) limit on any variable source of 0.66 \( \mu \)Jy, or an \( r \approx 24.4 \) source, which we take to be the limit of any supernova present at the time of our intermediate \( r \)-band observations (roughly 25 days postburst in the observer frame). Comparing this to redshifted lightcurves of SN 1998bw.

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**Table 2.** Gaussian component fitted to each energy band of the Swift light-curve (Figure 2). The peak is narrower and occurs earlier at harder energies.

| Bandpass (keV) | Gaussian centre (s) | Width (s) |
|---------------|---------------------|-----------|
| 0.3–10        | 234^{+16}_{-17}     | 185 ± 15  |
| 15–25         | 54 ± 9              | 39^{+10}_{-7} |
| 25–50         | 23^{+5}_{-4}        | 17^{+5}_{-4} |
| 50–100        | 8.6^{+1.5}_{-1.3}   | 6.6^{+1.8}_{-1.3} |
| 100–350       | 1.3^{+2.4}_{-1.0}   | 1.5^{+2.6}_{-0.9} |

**Figure 4.** The XRT-WT light-curves over 1.5–10 and 0.3–1.5 keV, together with the hardness ratio, showing an initial softening until \( \sim 140 \) s. After this time the hardness ratio remains approximately constant. The bottom panel shows the photon indices obtained from fitting the spectra with a single power-law.
(Galama et al. 1998), we determine that we would have observed a brightening due to a supernova component akin to SN 1998bw should it have been located at \( z \lesssim 0.7 \). This could be reduced if there were to be significant foreground extinction in the direction of GRB 080307. The moderately strong red R-K colour (R-K \( \sim 4 \)), suggests this may be the case, although there is no strong evidence for excess \( N_H \) from the X-ray data (see Section 4.3); thus, the lack of SN component does not place strong constraints on the redshift.

### 4.2 Possible constraints on the redshift

The lack of variation in the WHT data implies that those observations are dominated by the host galaxy. The absence of a supernova makes it unlikely that this burst occurred at low redshift [very high foreground extinction (\( A_V \sim 4.5 \) for a burst at \( z = 0.1 \)) would be required], despite the resemblance between this light-curve and that of the low-z GRB 060218 (Campana et al. 2006), whilst the magnitude of the host galaxy is entirely consistent with many other GRB hosts at \( z \lesssim 2 \) (see, e.g., Savaglio, Glazebrook & Le Borgne 2008).

Ukwatta et al. (2008) have investigated certain BAT parameters which may indicate when a burst is at a high redshift. Based on a limited sample of bursts, there are four criteria which together imply an 85% chance of a GRB having a redshift greater than 3.5; GRB 080307 fulfils these four points. Note, though, that the most distant GRB yet measured (GRB 080913; Greiner et al. 2008) did not satisfy all these criteria.

The XRT data from the second orbit onwards, when no spectral evolution is apparent and the light-curve is following a steady decline, gave an indication of an absorbing column in excess of the Galactic value of \( N_H \sim 1.6 \times 10^{21} \) cm\(^{-2} \) (at \( z = 0 \)); this is quite poorly constrained, however. This excess implies an upper limit on the redshift of \( z \lesssim 0.7 \) according to the method of Grupe et al. (2007a), though this increases to \( z \lesssim 3.5 \) using the lower limit on the excess column. This is in comparison to the lower limit estimate of \( z \gtrsim 3.5 \) as indicated by the BAT data (Section 4.3), although the detection of the host galaxy makes it unlikely that the burst occurred at such a redshift.

### 4.3 Spectral anomalies

If the additional column of \( N_H \sim 1.6 \times 10^{21} \) cm\(^{-2} \) is used when fitting time-sliced data through the hump, it is found that a single power-law is a poor fit, with a much softer component (either a second power-law, or a blackbody) being required as well. The ‘underlying’ power-law still shows the softening trend plotted in Figure 4 though.

GRB 060218 showed a similar rise in its early X-ray light-curve, and it was found that the spectra during this interval showed evidence for an expanding thermal component, corresponding to a supernova shock breakout (Campana et al. 2006). Fitting a similar model here, however, results in the temperature and radius of the blackbody component remaining approximately constant throughout (kT \( \sim 130–140 \) eV and an emitting radius of \( \sim 200D_{10kpc} \) km, where \( D_{10kpc} \) is the distance from the observer to the burst in units of 10 kpc), making it inconsistent with a shock breakout (a cooling temperature and expanding radius would be expected). No supernova has been identified for GRB 080307 either. It therefore seems unlikely that a thermal component is the explanation for the possible spectral curvature. A dual power-law model can equally well fit the spectrum, with photon indices of \( \Gamma \sim 3 \) and \( \sim 1 \) (both initially softening). Moretti et al. (2008) investigated additional spectral components in a sample of soft X-ray afterglow spectra. Their extra component was required to model an excess at higher energies, though, and there were no obvious humps in the light-curves, although there is some possible curvature in GRB 061110A (see http://www.swift.ac.uk/xrt_curves/00238109; Evans et al. 2007, 2009).

The excess absorbing column in GRB 080307 is not strongly significant, however, and no obvious explanation presents itself for multiple components. Thus, we simply mention this other unusual feature in passing and accept that there are residual uncertainties in the measurement of excess absorption.

### 4.4 Light-curve

The X-ray light-curve of GRB 080307 showed an unusual hump starting around 100 s after the burst. Below a number of possible explanations are considered.

#### 4.4.1 Flare

The XRT light-curve can be modelled as a single, unbroken power-law decay (with a large, superimposed hump), with no evidence for the series of breaks often found in the light-curves of X-ray afterglows (e.g. Nousek et al. 2006; Zhang et al. 2006). It is relatively unusual for a promptly-observed X-ray afterglow to follow a single decay, with only about 15% of Swift XRT light-curves showing no breaks before 100 ks (Evans et al. 2009).

Flares in X-ray afterglows have a typical relative width of \( \Delta(t)/t = 0.31 \pm 0.24 \) (Chincarini et al. 2007), where \( \Delta t \) is the Full Width at Half Maximum (FWHM = 2.3548\( \sigma \), where \( \sigma \), the Gaussian standard deviation width, is the number used in the Chincarini paper) and \( t \) is the time of the flare. The longest relative width in the Chincarini sample was \( \Delta t/t = 1.27 \) for GRB 051117A (Goad et al. 2007a),

![Figure 5. Ground-based optical data obtained using Gemini/GMOS (i-band), WHT/AUX (i and r-band) and FTS (R and i) and near IR (J, H and K-band) data from UKIRT.](http://www.swift.ac.uk/xrt_curves/00238109)
whereas GRB 080307 has $\Delta t/t \sim 1.48$. The rising emission at the start of the X-ray light-curve of GRB 080307 has a much longer duration than is typical for such an early time relative to the trigger. The other measurement of interest is the relative flux variability, $\Delta F/F$, which compares the flux at the peak of the ‘flare’ to that of the underlying power-law. In the case of GRB 080307, the values for $\Delta t/t$ and $\Delta F/F$ (1.48 and 6.17 respectively) lie in the patchy shells regime [see figure 1 in Ioka, Kobayashi & Zhang 2005 or figure 16 in Chincarini et al. (2007)], though could also be caused by refreshed shocks.

### 4.4.2 Reverse Shock

Kobayashi et al. (2007) discuss how synchrotron self-inverse Comptonisation (SSC) can cause electrons in the reverse shock region to be upscattered into the X-ray regime (reverse shocks are typically expected to radiate photons in the optical or infrared bands). They show that SSC emission can lead to an X-ray flare around the deceleration time using GRB 050406 (Romano et al. 2006) as an example. In such a case there needs to be a steep transition from the reverse shock to the forward one ($\sim t^{-(3p+1)/3}$ from figure 2 of Kobayashi et al.), but this decay phase could be very short, so can be accommodated by our data.

Reverse shock emission can only explain single flares and, in the thick shell case, is expected to overlap in time with the prompt emission. Figure 6 demonstrates that the X-ray emission in question does follow on from the BAT detection without any obvious gap, and thus reverse shock SSC emission is a possible origin of the hump seen in this burst.

### 4.4.3 Onset of afterglow

Instead of considering the initial X-ray rise as being a flare, however, it could be thought of as the onset of the afterglow - something which was expected to be seen in the pre-Swift era [e.g., Sari & Piran (1999a,b), Piro et al. (2005); Panaitescu & Vestrand (2008) and Molinari et al. (2007) also discuss rising optical afterglows]. The profile of the emission in GRB 080307 is quite smooth, rising gradually, something which is not generally the case for flares formed by internal shocks. Kobayashi & Zhang (2007) also discuss the onset of GRB afterglows, showing that a smooth bump can be produced by the forward shock emission.

Willingale et al. (2007) discuss a method of parametrizing BAT-XRT light-curves using one or two exponential-to-power-law functions to model the prompt and afterglow emission; see also O’Brien et al. (2006). If a single function is applied to the data, ignoring the broad hump as a flare, the start of the X-ray continuum emission is not well-fitted. Fitting one component to account for the prompt (BAT) emission and modelling the increase seen in the XRT as the onset of the afterglow produces a better fit, shown in Figure 6 [where the BAT data have been extrapolated into the 0.3-10 keV as described in O’Brien et al. (2006)]; the grey stars indicate data masked out when performing the fit.

The power-law starts to dominate the X-ray fit $\sim 480$ s after the trigger, with a decay slope of $\alpha \sim 2.12$. This is of the same order as the $\alpha \sim 1.95$ found in Section 3.2.1. There may be additional curvature during the later decay (after a few $10^4$ s), which could be related to deceleration of a thick fireball (Kobayashi & Zhang 2007 – see their figure 7; Lazzati & Begelman 2006).

Very few Swift X-ray light-curves show a rising component (other than flares), though most show a slow-decline, plateau phase (e.g. Nousek et al. 2006; Zhang et al. 2006). As shown by O’Brien et al. (2006) and Willingale et al. (2007), if the afterglow emission starts while the prompt emission is still bright, the initial rise may be hidden; then, when the prompt component fades sufficiently, the turn-over from the exponential to power-law decay is seen as a ‘flat-ish’ phase. If the XRT had started observing GRB 080307 even 50 s later, the rising portion would have been missed, and a (short) exponential-to-power-law decay is seen as a ‘flat-ish’ phase. If the XRT had started observing GRB 080307 even 50 s later, the rising portion would have been missed, and a (short) exponential-to-power-law decay would have been all that was seen at the start of the afterglow; this was the case for GRB 080810 (Page et al. in prep). Figure 6 shows that the function modelling the prompt, $\gamma$-ray emission drops very suddenly ($\alpha \sim 6$), which may be why the afterglow was prominently visible (that is, bright with respect to the prompt emission) from an early time. The sample of bursts in Willingale et al. (2007) contains only a small number of bursts with $\alpha_p > 5$. The BAT-XRT light-curve of GRB 060427 is plotted in that paper, demonstrating that, had it been a little brighter, the rise of the afterglow would probably have been briefly visible.

Following Molinari et al. (2007), the initial Lorentz factor, $\Gamma_0$, can be estimated from the peak time of the afterglow light-curve. Here we find that the X-ray light-curve peaks (assuming that this does correspond to the rise of the afterglow) around 234 s (Table 2); this corresponds to the deceleration time. Calculating the fluence over 1-1000 keV in the rest frame to be $2 \times 10^{-6}$ erg cm$^{-2}$, we find $\Gamma(t_{\text{peak}}) \sim 156(\eta_0 n_0)^{-1/2}(1+z)^{3/8}$, where $\eta_0$ is the radiative efficiency in units of 0.2 and $n_0$ is the particle density of the (homogeneous) surrounding medium in cm$^{-3}$. Since $\Gamma(t_{\text{dec}})$, which is equal to $\Gamma(t_{\text{peak}})$ here, is expected to be half of the initial Lorentz factor, $\Gamma_0 = 312(\eta_0 n_0)^{-1/2}(1+z)^{3/8}$.
For the mean Swift redshift of 2.26 (Jakobsson et al. 2006),
this corresponds to a value of 486 (assuming typical values of 
$\eta = 0.2$ and $n = 1 \text{ cm}^{-3}$ from Bloom, Frail & Kulkarni 2003).
The limit on the redshift from the Grupe et al. $N_H$
analysis suggests an upper limit of $\Gamma_0 < 562$.

Swift-XRT light-curves do not typically contain
smoothly rising features, although flares are seen in about
50 per cent of bursts (e.g., Chincarini et al. 2007; Fal-
cone et al. 2007). Utilising the light-curve repository (Evans
et al. 2007, 2009), we find there are a small number
of bursts where a possible early rise is seen which does
not look like a standard flare (e.g. GRB 070328 – Mark-
wardt et al. 2007; GRB 070714B – Racusin, Barbier & Landsman 2007; GRB 080205 – Markwardt et al. 2008;
GRB 080229A – Cannizzo et al. 2008; GRB 080721 – Mar-
shall et al. 2008). However, in these cases the light-curve can
be sufficiently well-modelled with a series of abruptly bro-
ken power-laws, rather than requiring a smooth turnover.
There are also bursts which show a more gradual roll-over
than can be well-modelled by a series of broken power-
laws (GRB 070508 – Grupe et al. 2007b; GRB 080319C
– Starling et al. in prep; GRB 080503 – Mao et al. 2008;
GRB 080523 – Stroh et al. 2008; GRB 080605 – Sbaru-
fatti et al. 2008). More recently, GRB 081028 (Guidorzi,
Margutti & Mao 2008; Schady & Guidorzi 2008) showed
smoothly rising emission, but in this case beginning around
10 ks after the trigger time, and reaching a peak about 10 ks later. For this burst, no spectral softening is ap-
parent (see http://www.swift.ac.uk/xrt_curves/00332851).
It remains unclear whether these bursts are displaying the
same phenomenon as GRB 080307 or are unrelated, how-
ever.

5 SUMMARY

GRB 080307 showed a long, smooth hump early in its X-ray
afterglow light-curve. Although superficially similar to the
supernova shock breakout of GRB 060218, in this case no
evidence for radial expansion was seen. The hump is unique
in the Swift dataset, although may possibly be less well ob-
erved in a few other bursts. We have considered various ori-
gins for the hump: a flare due to a patchy shell or a refreshed
shock, reverse shock SSC, and the onset of the afterglow.

The long duration of the hump would make this a very
slow flare for early times and the spectral softening seen
at the start of the curvature is also unusual for a flare. In
the case of the reverse shock interpretation, the transition
to the forward shock would have to be very short to be
hidden within the data. All of these mechanisms are viable
explanations of the observed properties, however the onset
of the X-ray afterglow remains a natural interpretation.

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