Magnetic field geometry and magnetospheric environment of the strongly magnetic Of?p star NGC 1624-2

S. P. Järvinen1*, S. Hubrig1, M. Schöller2, M. Küker1, I. Ilyin1, S. D. Chojnowski3

1Leibniz-Institut für Astrophysik Potsdam (AIP), An der Sternwarte 16, 14482 Potsdam, Germany
2European Southern Observatory, Karl-Schwarzschild-Str. 2, 85748 Garching, Germany
3Apache Point Observatory and New Mexico State University, P.O. Box 59, Sunspot, NM, 88349-0059, USA

Accepted XXX. Received YYY; in original form ZZZ

ABSTRACT
NGC 1624-2 is an O7If?p star with a reported probable polar magnetic field strength \( \geq 20 \) kG, which is the strongest magnetic field ever measured in an O-type star. We study the variability of the mean longitudinal magnetic field \( \langle B_z \rangle \) and the mean field modulus \( \langle B \rangle \) to obtain constraints on its field geometry. Only one magnetic pole is observable over the rotation cycle. The approximately sinusoidal variation of \( \langle B_z \rangle \) and the ratio of the values of the extrema of \( \langle B \rangle \) indicate that there is an important component of the field that is dipolar. The \( \langle B_z \rangle \) values measured over the rotation cycle are in the range from \(-0.2\) to \(4.5\) kG, whereas the values for \( \langle B \rangle \) vary between \(9\) and \(12\) kG. The \( \langle B_z \rangle \) values obtained using the \( \text{O} \, \text{m} \) \( \lambda \)7455 emission line are in the range from \(0.4\) to \(2.3\) kG and show a variability pattern similar to that detected for the absorption lines. The fact that the phase of the \( \langle B_z \rangle \) minimum coincides with the phase of the \( \langle B \rangle \) maximum, indicates that the field structure must significantly depart from a centred dipole. Further, we discuss the nature of the observed variable Stokes \( V \) profiles corresponding to a longitudinal field of negative polarity detected in the emission \( \text{He} \, \text{i} \) lines and present the first MHD numerical simulations of the gas flow in the magnetosphere of this star.

Key words: stars: early-type – stars: individual: NGC 1624-2 – stars: magnetic fields – stars: winds, outflows – polarization – MHD

1 INTRODUCTION
Recent observations indicate that probably only a small fraction of about 7 ± 3% of O-type stars with masses exceeding 18 \( \text{M}_\odot \) (Grunhut et al. 2017) and about 6 ± 3% of early B- and O-type stars (Schöller et al. 2017) have measurable, mostly dipolar magnetic fields. Among the polarimetrically studied O-type stars, all five Galactic stars with Of?p classification possess measurable magnetic fields. These stars exhibit recurrent, and apparently periodic, spectral variations in Balmer, \( \text{H} \, \text{e} \), \( \text{C} \, \text{m} \), and \( \text{Si} \, \text{m} \) lines, sharp emission or P Cygni profiles in \( \text{He} \, \text{i} \) and the Balmer lines, and a strong \( \text{C} \, \text{m} \) blend in emission around \( 4650 \) Å (Walborn 1972). The presence of variable emission in the \( \text{C} \, \text{m} \) blend and in \( \text{H} \, \text{o} \) in Of?p stars is indicative of circumstellar structure, related to their magnetospheres. Previous measurements of the mean longitudinal magnetic field for the Of?p star NGC 1624-2 carried out with ESPaDOnS (Echelle SpectroPolarimetric Device for the Observation of Stars) at the Canada-France-Hawaii Telescope between 2012 February 1 and 9 indicated \( \langle B_z \rangle = 5.35 \pm 0.5 \) kG with a polar magnetic field strength of about \(20\) kG, which is the strongest magnetic field ever measured in an O-type star (Wade et al. 2012a). All the other Of?p stars with studied magnetic field geometries have \( B_{\text{pole}} \leq 2.6 \) kG (David-Uraz et al. 2019). The existence of a giant wind-fed dynamo magnetic field around NGC 1624-2 was discussed by Petit et al. (2015). A very recent study by Kurtz et al. (2020) reported the detection of coherent pulsation modes in NGC 1624-2 based on Transiting Exoplanet Survey Satellite (TESS; Ricker et al. 2014) high-cadence photometry.

The study of spectral variability using archival optical spectroscopic observations indicated \( v \sin i \leq 3 \) km s\(^{-1}\) and a rotation period of \(157.99 \pm 0.94\) d (Wade et al. 2012a). Due to the presence of the strong magnetic field, slow rotation, and the low projected rotational velocity, a few spectral lines showed magnetic splitting into Zeeman components corresponding to a maximum mean magnetic field modulus of about \( \langle B \rangle = 14 \pm 1\) kG. Furthermore, Wade et al. (2012a) reported the detection of reversed Stokes \( V \) profiles associated with weak, high-excitation emission \( \text{O} \, \text{m} \) lines, and suggested that they may form in the low-velocity plasma confined in closed magnetic loops above the stellar surface. They measured a longitudinal field \( \langle B_z \rangle = 2.58 \pm 0.7 \) kG, which is by more than a factor of 2 smaller than the field measured using photospheric absorption lines. On the other hand, the emission lines \( \text{He} \, \text{i} \) \( 5876, 6678, \) and 7065 exhibited Stokes \( V \) profiles with signs consistent with those of the photospheric absorption lines, but yielding a negative mean longitudinal magnetic field. Unfortunately, no field measurements using these lines were presented in the work of Wade et al. (2012a).
The He I emission lines are marked in the paper of Wade et al. as “wind lines”. This expression was already used by Howarth et al. (2007) for Hα, where the authors discuss some lines as contaminated by “windy” emission. Several studies of magnetic O-type stars concluded that these lines are not formed in the wind, but are possibly resulting from confined circumstellar material (e.g. Grunhut et al. 2009). In the context of the magnetically confined wind-shock model, it was suggested that most emission comes from scattering by the cooling disk and very little from the wind (Donati et al. 2002).

The detected remarkable characteristics of NGC 1624-2 make it an excellent laboratory to study the impact of a very strong magnetic field on the behaviour of different elements in the stellar photosphere and circumstellar environment. In this work, we discuss for the first time the magnetic field geometry of NGC 1624-2 based on the variability of the longitudinal magnetic field measured using photospheric lines, the longitudinal field measurements using the emission O iii and He I lines, and the variability of the magnetic field modulus over the rotation period. Finally, using the NIRVANA code, we discuss the first magnetohydrodynamical (MHD) numerical simulations carried out to understand the field geometry, mass distribution, and the gas motions in the vicinity of NGC 1624-2.

2 OBSERVATIONS

A few polarimetric observations were obtained at the beginning of 2012 (Wade et al. 2012a) with the ESPaDOnS spectropolarimeter installed at the 3.6 m Canada-France-Hawaii Telescope (CFHT) together with one observation from its twin Narval, installed at the 2 m Bernard Lyot Telescope (BLT) on Pic-du-Midi. Six additional ESPaDOnS observations from 2012 September to 2013 January were presented by Grunhut et al. (2017). Subsequently, sixteen ESPaDOnS observations were obtained between 2013 August and 2015 September. They are described in more detail in David-Uraz et al. (2020). All the ESPaDOnS data are publicly available in the CFHT Science Archive1 and both ESPaDOnS and Narval data can be retrieved from the PolarBase2 archive (Petit et al. 2014).

Additionally, one Potsdam Echelle Polarimetric and Spectroscopic Instrument (PEPSI; Strassmeier et al. 2015, 2018) spectrum has been recorded in linear polarized light in 2017 October at the 2x8.4 m Large Binocular Telescope (LBT) on Mt. Graham, Arizona and one unpolarized spectrum was also obtained with the Astrophysical Research Consortium Echelle Spectrograph (ARCES; Wang et al. 2003) on the ARC 3.5 m telescope at the Apache Point Observatory (APO) on 2019 March 1.

The logbook of all available observations is presented in Table 1. The recorded linear polarization PEPSI Stokes Q and U spectra and the ARCES spectrum were however very noisy, hence only the Stokes I spectrum was used for measuring the magnetic field modulus. The resolving power of all archival observations is \( \sim 65000 \), while the spectral resolution of the PEPSI spectrum is 130000 whereas the ARCES spectrum has a resolution of only 31 500. As most of the spectra have moderate signal-to-noise ratios (S/N) in the Stokes I spectra, mostly in the range of 100–150, some spectra obtained during the same night or, in one case, during subsequent nights have been combined to increase the signal for the Stokes V profiles. However, in cases where the S/N of the spectrum was significantly lower than that of the other spectra obtained during the same night, only the spectra with better S/N were used for the analysis of the mean longitudinal magnetic field. Some spectra with low S/N have been included in the analysis of the line intensity variability or used for the measurements of the mean magnetic field modulus, if the Zeeman splitting was detectable on them. Further, to increase the accuracy of the mean longitudinal magnetic field determination, we applied the least-squares deconvolution technique (LSD) described in detail by Donati et al. (1997). To create line masks for the measurements, we have used the Vienna Atomic Line Database (VALD; e.g., Kupka et al. 2011; Ryabchikova et al. 2015), adopting the stellar parameters of NGC 1624-2, \( Teff = 35000 \) K and \( \log g = 4.0 \) (Wade et al. 2012a). The mask for the photospheric absorption lines includes spectral lines belonging to the elements He, H, C, N, O, Ne, and Mg.

| Configuration    | HD      | Exp. time [s] | S/N |
|------------------|---------|---------------|-----|
| CFHT + ESPaDOnS  | 55958.719 | 2400 | 109 |
| CFHT + ESPaDOnS  | 55959.719 | 2400 | 88  |
| CFHT + ESPaDOnS  | 55960.717 | 2400 | 71  |
| CFHT + ESPaDOnS  | 55961.717 | 2400 | 112 |
| CFHT + ESPaDOnS  | 55966.725 | 2400 | 90  |
| TBL + Narval     | 56011.330 | 4800 | 78  |
| CFHT + ESPaDOnS  | 56197.950− | 5400 | 146 |
| CFHT + ESPaDOnS  | 56198.014− | 5400 | 146 |
| CFHT + ESPaDOnS  | 56200.974 | 5400 | 130 |
| CFHT + ESPaDOnS  | 56201.039 | 5400 | 148 |
| CFHT + ESPaDOnS  | 56270.751− | 5400 | 115 |
| CFHT + ESPaDOnS  | 56270.816− | 5400 | 138 |
| CFHT + ESPaDOnS  | 56285.761− | 5400 | 93  |
| CFHT + ESPaDOnS  | 56285.826 | 5400 | 127 |
| CFHT + ESPaDOnS  | 56285.890 | 5400 | 131 |
| CFHT + ESPaDOnS  | 56293.733− | 5400 | 136 |
| CFHT + ESPaDOnS  | 56293.798− | 5400 | 138 |
| CFHT + ESPaDOnS  | 56295.815− | 5400 | 19  |
| CFHT + ESPaDOnS  | 56352.126 | 5160 | 156 |
| CFHT + ESPaDOnS  | 56354.063 | 5160 | 154 |
| CFHT + ESPaDOnS  | 56354.124 | 5160 | 135 |
| CFHT + ESPaDOnS  | 56549.039− | 5160 | 146 |
| CFHT + ESPaDOnS  | 56549.102− | 5160 | 141 |
| CFHT + ESPaDOnS  | 56561.028 | 5160 | 150 |
| CFHT + ESPaDOnS  | 56561.091 | 5160 | 123 |
| CFHT + ESPaDOnS  | 56613.913− | 5160 | 112 |
| CFHT + ESPaDOnS  | 56614.045− | 5160 | 125 |
| CFHT + ESPaDOnS  | 56621.039− | 5160 | 140 |
| CFHT + ESPaDOnS  | 56621.102 | 5160 | 135 |
| CFHT + ESPaDOnS  | 56665.828− | 5160 | 123 |
| CFHT + ESPaDOnS  | 56665.889− | 5160 | 149 |
| CFHT + ESPaDOnS  | 57289.978− | 5160 | 82  |
| CFHT + ESPaDOnS  | 57290.041 | 5160 | 106 |
| CFHT + ESPaDOnS  | 57291.059 | 5160 | 108 |
| LBT + PEPSI      | 58042.022 | 3600 | 70  |
| ARC + ARCES      | 58543.653 | 3600 | 98  |

1 https://www.cadc-cda.hia-tha.nrc-cnrc.gc.ca/en/cfht/
2 http://polarbase.irap.omp.eu/
The mean longitudinal magnetic field \( \langle B_z \rangle \) is the velocity shift from the line centre, in km s\(^{-1} \), and \( \lambda_0 \) and \( \beta_0 \) are the normalization values of the line weights. We have adapted for them here the average values of, respectively, the wavelengths (in nm) and the effective Landé factors of all the lines used to compute the LSD profile. The results of our magnetic field measurements using the line mask for the photospheric absorption lines, including the Si lines, the line mask without the Si lines, and the measurements using the high excitation emission line O\( \text{iii} \) \( \lambda \)74555 and the strongest emission lines He\( \text{i} \) \( \lambda \)5876 and He\( \text{i} \) \( \lambda \)6678 are presented in Fig. 1 and in Table 2 in Columns 3 to 7. All LSD profiles for the selected line masks, as well as Stokes \( I \), \( V \), and null profiles for the individual emission lines are shown in Figs. A1–A6 in the Appendix. The average effective Landé factors used for the normalization for the line mask including the Si lines and the mask without the Si lines are 1.55 and 1.24, respectively. The corresponding wavelengths are 5013 and 5224 Å. For the profiles calculated using all lines and the line list without the Si lines an integration was made from \(-120\) to \(+25\) km s\(^{-1} \), for O\( \text{iii} \) \( \lambda \)7455 from \(-70\) to \(+10\) km s\(^{-1} \), for He\( \text{i} \) \( \lambda \)5876 and for He\( \text{i} \) \( \lambda \)6678 from \(-200\) to \(+150\) km s\(^{-1} \).

In most cases the false alarm probability (FAP) for the field detections is less than \( 10^{-6} \). According to Donati, Semel, & Rees (1992), a Zeeman profile with FAP \( \leq 10^{-6} \) is considered as a definite detection, \( 10^{-5} < \text{FAP} \leq 10^{-3} \) as a marginal detection, and FAP \( > 10^{-3} \) as a non-detection. Our measurements show that the \( \langle B_z \rangle \) values become lower if the Si\( \text{iv} \) lines are included in the line mask. A possible scenario to explain these measurements would be an inhomogeneous surface element distribution with a region of Si concentration located closer to the magnetic equator. It is also possible that we observe a filling of the lines by emission. Our \( \langle B_z \rangle \) measurements in the range from \(-0.2\) to \(+4.5\) kG obtained for the sample of the photospheric absorption lines are in good agreement with the \( \langle B_z \rangle \) values reported by Grunhut et al. (2017), but disagree with the measurement of \( \langle B_z \rangle = 5.35 \) kG reported by Wade et al. (2012a). Assuming a limb-darkening coefficient \( u = 0.3 \) (Claret & Bloemen 2011) and a centred dipolar magnetic field, we estimate the dipole strength of NGC 1624-2 \( B_0 = 16 \) kG according to Equation (1) of Preston (1967). The dashed line in Fig. 1 represents the best sinusoidal least-squares fit solution for the phase curve defined by our magnetic field measurements using the line mask for photospheric absorption lines, without the Si\( \text{iv} \) lines. The rotation phase presented in Table 2 in the second column is calculated using the rotation period of 157.99 d adopted by Wade et al. (2012a) and the initial epoch for the phase origin HJD\( _0 \) = 2455955.94 ± 0.21 corresponding to the maximum of the longitudinal field phase curve. The fitted sinusoid has an amplitude \( A(B_z) = 2132 \pm 364 \) G and a mean value \( \langle B_z \rangle = 2063 \pm 328 \) G, thus \( \langle B_z \rangle_{\text{max}} = 4195 \) G and \( \langle B_z \rangle_{\text{min}} = -69 \) G.

Knowing the rotation period, we can estimate the equatorial velocity and the inclination of the rotation axis of the star \( i \). Assuming a radius \( R = 9.7 R_\odot \) (Petit et al. 2013) and a projected rotation velocity \( v \sin i \leq 3 \) km s\(^{-1} \) (Wade et al. 2012a), we obtain \( i \leq 75^\circ \). Using the well known relations developed by Stibbs (1950) and Preston (1967) for a centred magnetic dipole tilted to the rotation axis by angle \( \beta \), we estimate an obliquity angle \( \beta \geq 15.5^\circ \). However, with \( v \sin i = 1 \) km s\(^{-1} \), we obtain \( i = 19^\circ \) and \( \beta = 72^\circ \), i.e. nearly a flip of the two axes, showing that the exact geometry of the magnetic field can not be determined with the still too loose constraint on \( v \sin i \). The measured values of the longitudinal magnetic field of NGC 1624-2 with predominantly positive polarity imply that only one magnetic pole is observable over the rotation cycle. The dynamic spectra of emission and photospheric lines show a single-wave variation of their intensity: the intensity of the lines H\( \alpha \), He\( \text{ii} \) \( \lambda \)5846, He\( \text{i} \) \( \lambda \)5876 and He\( \text{i} \) \( \lambda \)6678 are the strongest to the positive extremum of \( \langle B_z \rangle \) whereas the photospheric absorption lines of He\( \text{i} \), C\( \text{iv} \), Si\( \text{iv} \), and the emission O\( \text{iii} \) lines are the strongest close to the rotation phase 0.5. Examples of the variability of different lines are presented in Fig. 2.

### 3 MAGNETIC FIELD MEASUREMENTS

#### 3.1 Mean longitudinal magnetic field \( \langle B_z \rangle \)

The mean longitudinal magnetic field is determined by computing the first-order moment of the LSD Stokes \( V \) profile according to Mathys (1989):

\[
\langle B_z \rangle = -2.14 \times 10^{11} \frac{\int \nu V(\nu) d\nu}{\lambda_0 \beta_0 \int [I_c - I(\nu)] d\nu}.
\]  

where \( \nu \) is the velocity shift from the line centre, in km s\(^{-1} \), and \( \lambda_0 \) and \( \beta_0 \) are the normalization values of the line weights. We have adapted for them here the average values of, respectively, the wavelengths (in nm) and the effective Landé factors of all the lines used to compute the LSD profile. The results of our magnetic field measurements using the line mask for the photospheric absorption lines, including the Si lines, the line mask without the Si lines, and the measurements using the high excitation emission line O\( \text{iii} \) \( \lambda \)74555 and the strongest emission lines He\( \text{i} \) \( \lambda \)5876 and He\( \text{i} \) \( \lambda \)6678 are presented in Fig. 1 and in Table 2 in Columns 3 to 7. All LSD profiles for the selected line masks, as well as Stokes \( I \), \( V \), and null profiles for the individual emission lines are shown in Figs. A1–A6 in the Appendix. The average effective Landé factors used for the normalization for the line mask including the Si lines and the mask without the Si lines are 1.55 and 1.24, respectively. The
the magnetic field vector, weighted by the local line intensity, are presented in Table 3 and Fig. 3. The field modulus was calculated by fitting the magnetically resolved C iv λ5812 line with two Gaussian profiles and using the relation given e.g. by Hubrig & Nesvacil (2007). The uncertainties in the presented ⟨B⟩ values are due to the accuracy of the measured wavelengths of the blue and red split line components. Our measurements indicate a field intensity between 9 and 12 kG differ from the measurements reported by Wade et al. (2012), who reported ⟨B⟩ = 15 kG for the λ5801 line and ⟨B⟩ = 13 kG for the λ5812 line using the Narval spectrum. This difference can likely be explained by the Zeeman pattern of the C iv λ5812 line. We interpret in our measurements the splitting as a pseudo-doublet, which underestimates the strength of the field modulus.

The approximately sinusoidal variation of ⟨B⟩ and the ratio of the values of the extrema of ⟨B⟩ indicates that there is an important component of the field that is dipolar. Further, a comparison of the variation curves of the mean magnetic field modulus and the mean longitudinal magnetic field shows that the minimum of the mean magnetic field modulus corresponds to the maximum of the mean longitudinal magnetic field. This indicates that the field structure must significantly depart from a centred dipole. It is possible that we see a small shift between the modulus minimum with respect to the positive extremum of the ⟨B⟩, although the fidelity of such a shift is only barely supported, given the rather large dispersion of the data points. A future confirmation of the presence of a shift is necessary, as this would indicate that the magnetic field structure of NGC 1624-2 slightly departs from a simple dipole.

### 3.3 Emission lines

We confirm the conclusion of Wade et al. (2012a) that the ⟨B⟩ values measured using the high-excitation O iii emission lines, suggested to form in the low-velocity plasma confined in closed magnetic loops above the stellar surface, are significantly lower than those obtained for photospheric absorption lines. The ⟨B⟩ values obtained using the O iii λ4545 line are in the range from 0.4 to
Table 3. The mean magnetic field modulus ($B$) values obtained in the spectra with measurable splitting in the C iv $\lambda$5812 line.

| Date     | Phase | $\langle B \rangle$ (G) |
|----------|-------|-------------------------|
| 55959.719 | 0.024 | 9731±431                |
| 55960.717 | 0.030 | 9433±526                |
| 55961.717 | 0.037 | 9650±950                |
| 55966.725 | 0.068 | 10927±300               |
| 56011.330 | 0.351 | 11546±1053              |
| 56197.950 | 0.532 | 11424±304               |
| 56198.014 | 0.532 | 10975±409               |
| 56200.974 | 0.551 | 11059±356               |
| 56201.039 | 0.551 | 10990±360               |
| 56270.751 | 0.993 | 10299±385               |
| 56270.816 | 0.993 | 10973±491               |
| 56285.826 | 0.088 | 9253±714                |
| 56285.890 | 0.088 | 9682±423                |
| 56293.733 | 0.138 | 10322±616               |
| 56293.799 | 0.138 | 9932±299                |
| 56352.126 | 0.647 | 11099±301               |
| 56354.063 | 0.659 | 11384±310               |
| 56354.124 | 0.660 | 11405±377               |
| 56549.039 | 0.754 | 11586±497               |
| 56549.102 | 0.754 | 11372±408               |
| 56561.028 | 0.830 | 10616±379               |
| 56613.913 | 0.165 | 10751±485               |
| 56614.045 | 0.165 | 9127±950                |
| 56621.039 | 0.210 | 10433±713               |
| 56621.102 | 0.210 | 9510±475                |
| 56665.828 | 0.493 | 10840±341               |
| 56665.889 | 0.494 | 10803±346               |
| 57289.978 | 0.444 | 11396±319               |
| 57290.041 | 0.444 | 10886±294               |
| 57291.059 | 0.451 | 11683±555               |
| 58042.022 | 0.204 | 11122±600               |
| 58543.653 | 0.379 | 10913±425               |

Figure 3. Mean magnetic field modulus ($B$) of NGC 1624-2 measured using the split components of the magnetically resolved line C iv $\lambda$5812 against the rotation phase.

Figure 4. As Fig. 1, but $\langle B_z \rangle$ measured using the He i 5876 Å and He i 6678 Å lines.

2.3 kG and show a variability pattern similar to that detected for the absorption lines.

Interestingly, we detect in the spectra the forbidden [O i] 6300 and 6363 lines, indicating the presence of low-density atomic material in the vicinity of NGC 1624-2. The formation mechanism of these lines is unclear. It was suggested that they originate from disk winds or regions where stellar UV radiation impinges on the disk surface in pre-main sequence Herbig Ae/Be stars (e.g. Finkenzeller 1985; Corcoran & Ray 1997). It is also possible that the forbidden [O i] 6300 and 6363 lines are formed in magnetospheric accretion columns (Muzerolle et al. 2004). The variability of the strength of these lines is presented in Fig. B1. They are the strongest at the rotational phases 0.03 and 0.35 and do not show rotational modulation. This suggests that they are most likely formed beyond the Alfven radius.

Our analysis of the observations of NGC 1624-2, in the scenario of an oblique magnetic rotator, indicates that stronger emissions in the hydrogen and helium emission lines are detected when the magnetically-confined cooling disk is seen closer to face-on, while the emission contribution is reduced when the cooling disk is seen closer to edge-on (see Fig. 2).

No measurements using the He i emission lines were presented in the study of Wade et al. (2012a). Our ($B_z$) values are predominantly negative and, similar to the measurements of other spectral lines, show clear variability over the rotation period, with the negative ($B_z$) extremum observed around phase 0, where the photo-spheric absorption lines show a positive ($B_z$) extremum. The distribution of our measurement values over the rotation period obtained using the He i $\lambda$5876 and $\lambda$6678 lines, which are the strongest in the $2^3P1^n^1D$ and $2^1P1^n^1D$ series, respectively, is presented in Fig. 4. The significant data point dispersion and the presence of two outliers in the measurements of He i $\lambda$5876 close to the rotation phase 0.2 is caused by the short-term variability of the line intensity on a time scale of several days, and is likely a consequence of either the pulsational variability recently detected by Kurtz et al. (2020) or a time-variable structure of gas flows within the magnetosphere (ud-Doula et al. 2013).

The origin of the emission in the He i lines in the red spectral region is not well understood. A long time ago, Auer & Mihalas (1972) showed that NLTE effects are quite small for He i lines in the blue-violet region of the spectrum, whereas the red He i lines, in particular He i $\lambda$5876 and He i $\lambda$6678, are strongly affected by departures from LTE. A large discrepancy between observations and their NLTE results occurred for He i $\lambda$5876, where the observed lines are consistently stronger than computed. However, the presented NLTE calculations always remained in absorption. Only at the highest temperatures of about 50,000 K the NLTE results predicted that lines come into emission.
Several mechanisms were suggested since then to understand the appearance of the emission observed in these lines. To explain the emissions in the He\textsc{i}\ 6587,6\,6678 lines in the early type star \lambda\ Eri, Smith et al. (1997) suggested that emission can be produced in a dense slab that is optically thick in the Lyman continuum. To produce a large enough column density for an observable feature, the requirement was made that the slab must be cohesive and is confined in a magnetic loop to resist the acceleration of the ambient wind.

The presence of variable Zeeman features corresponding to a longitudinal field of negative polarity observed in the He\textsc{i} emission lines, while high-excitation emission O\textsc{i} in lines show the same polarity as the photospheric absorption lines, is puzzling and currently difficult to interpret. One alternative hypothesis could be that the He\textsc{i} Stokes \textit{I} and \textit{V} profiles have a composite structure and present a contribution of several components. In Fig. 5 we present the Stokes \textit{I}, \textit{V}, and null profiles for all three He\textsc{i} emission lines at two rotational phases. Indeed, depending on the line, the observed Stokes \textit{I} profiles indicate a rather complex variable structure consisting of at least two or three components (see also Figs. A4–A6). It is quite possible that there is a correspondence between the different parts of the Stokes \textit{I} profiles and those of the Stokes \textit{V} profiles. The resulting composite Stokes \textit{V} profiles would then present just a blend of the individual Stokes \textit{V} profiles corresponding to different parts of the Stokes \textit{I} profiles. It is noteworthy that for the majority of the obtained observations, the shape of the Stokes \textit{V} profiles displayed in Figs. A4–A6 cannot be explained in terms of the Zeeman effect: if these profiles are indeed formed in a hydrostatic stellar atmosphere lacking gradients in velocity and magnetic field as a function of optical depth (e.g. Sanchez Almeida, Collados & del Toro Iniesta 1989; Lópe Ariste 2002), they always have the zero-order moment equal to zero, that is the integral of Stokes \textit{V} over the region of the spectral line in Stokes \textit{I} is equal to zero. The observed differences between the areas corresponding to the blue and red parts of the Stokes \textit{V} profiles calculated for all three He\textsc{i} emission lines are presented in Fig. 6. The differences were calculated only in rotation phases with a lower impact of the noise on the shape of the Stokes \textit{V} profiles. As this figure shows, for all three lines the differences between the areas corresponding to the blue and red parts of the Stokes \textit{V} profiles are significant.

The changes in emission strength of the He\textsc{i} lines and the different sign of the longitudinal magnetic field in these lines can also be understood within the following scenario. For a dipole magnetic field, the magnetic field lines reverse their orientation at the equator relative to the pole. If emitting material accumulates in the magnetosphere at the magnetic equator, it experiences a magnetic field of opposite orientation to that at the magnetic pole. The longitudinal magnetic field measured in the He\textsc{i} lines as seen in emission from the magnetospheric material at the magnetic equator would have a sign opposite from the longitudinal field measured from a photospheric absorption line. The extended region near the equator with an elevated density seems like a plausible source of circumstellar emission (see Fig. 7 for an illustration). The magnetic field there has an opposite orientation to that dominating the upper hemisphere of the photosphere. Then, as the star rotates, the longitudinal component of the magnetic field in the emitting region becomes weaker, and the projected area of the emitting region diminishes, and possibly becomes partially occulted by the star.
Another alternative hypothesis could be related to the complex structure of the magnetospheric circumstellar environment of NGC 1624-2 and the possible presence of a reconnection-driven emission process. To understand the behaviour of the gas flow and the mass distribution within the magnetosphere, we carried out first MHD numerical simulations for this star.

4 MHD NUMERICAL SIMULATIONS

A dipolar magnetic field with a polar field strength of 16 kG will strongly alter the gas flow from the star and inhibit mass loss. The impact of the magnetic field is described by the wind confinement parameter, \( \eta = B^2_{\text{eq}}R^2_*/Mv_\infty \), where \( B_{\text{eq}} \) is the field strength on the stellar surface at the magnetic equator, \( R_* \) the stellar radius, \( M \) the mass loss rate with no magnetic field, and \( v_\infty \) the terminal velocity of the stellar wind (ud-Doula & Owocki 2002). For NGC 1624-2, \( \eta_1 \times 10^4 \) was found by Wade et al. (2012a), which results in a value of 11.3 \( R_* \) for the Alfvén radius, which indicates the extent of the magnetically-dominated region. Within that region, the magnetic field will keep its dipole geometry and gas flowing from the star at low to mid-latitudes will be trapped. Only gas originating from the polar caps leaves the system unimpeded along open field lines.

To study the field geometry, mass distribution, and gas motions in the vicinity of the star, we ran a series of numerical simulations. The general model setup is described in Küker (2017). We used the NIRVANA MHD code (Ziegler 2008) in spherical polar coordinates. The geometry is axisymmetric with respect to the polar axis of the magnetic field and the gas is assumed isothermal, with the gas temperature being the effective temperature of the star. No rotation is included. As the model does not include cooling, magnetic diffusion has been added with a diffusion coefficient of 10\(^{15}\) cm\(^2\) s\(^{-1}\). The stellar and wind parameters have been adopted from Wade et al. (2012a). We thus use a stellar mass of 34 solar masses, a stellar radius of 10 solar radii and an effective temperature of 35 000 K. The mass loss rate is 1.6 \times 10^{-7} M_{\odot} yr\(^{-1}\) and the terminal wind velocity is 2785 km s\(^{-1}\). The start solution is a dipole magnetic field and a velocity law given by

\[
\nu(r) = v_\infty \left( 1 - \frac{R_*}{R} \right).
\]

The initial mass distribution is determined by the velocity law (2) and the mass loss rate, i.e.

\[
\rho = \frac{M}{4\pi \nu r^2}.
\]

A radiative force term after Castor, Abbott & Klein (1975) with \( \dot{Q} = 550 \) and \( \alpha = 0.55 \) is used to drive the wind. The parameters have been chosen such that without the magnetic field a stationary solution with the correct wind parameters is produced. We then repeat the computation with a dipole field included. To limit the cost in CPU time, the outer boundary of the simulation box is placed at 5 solar radii and a polar field strength of 4 kG is chosen. While this is smaller than the observed field strength of 16 kG, it still leads to an extended region of magnetic confinement with \( \eta_1 \approx 700 \) and an Alfvén radius of 5.4 stellar radii. Within the simulation box, the magnetic field geometry stays close to the original dipolar geometry and only a fraction of the gas flowing from the stellar surface can escape the system. The mass loss rate found by Wade et al. (2012a) is based on a prescription by Vink, de Koter & Lamers (2001) that does not take the magnetic field into account. It therefore refers to the amount of gas that would leave the system if there was no magnetic field.

In the initial phase of the simulation run, the gas flow and mass density undergo drastic changes. While the gas density at high latitude drops, gas piles up in the equatorial plane, where a thin disk forms from which filaments grow that are aligned with the magnetic field. Figure 7 shows a snapshot of the mass density taken at a time when the system has reached a semi-stationary state insofar as the general mass distribution with a high mass concentration at low to mid latitudes does not evolve much any more. The shape and length of the individual filaments is highly variable, though, and the system does not reach a truly stationary state. Note that all the field lines shown are closed. Compared to the original dipole, the field is slightly compressed by the gas load but the general field topology is unchanged. This is generally different for weak surface fields or at greater distances from the star, where the field lines are stretched away from the star and a thin disk of outflowing material forms.

Figure 8 shows a snapshot from a simulation with a weaker field of 1 kG. The magnetosphere is much less extended both in radius and latitude and far away from the star the magnetic field is stretched out into a configuration resembling a split monopole with a current sheet in the equatorial plane. The gas forms a thin disk of enhanced mass density in the equatorial plane. The disk is fragmented and the magnetic field is frozen into the fragments. In the equatorial plane the magnetic field forms an elongated tail. At the tip, reconnection occurs and magnetic flux is carried away by the disk material. We expect a similar picture in case of stronger fields, but at distances from the star outside our current simulation box. Note also that for the stronger magnetic field in Fig. 7, the regions of the stellar surface where open field lines originate is limited to a small range of latitudes around the poles. We expect this effect to be even stronger with a dipole field strength as large as 20 kG. The right panel of Fig. 7 shows a zoom-in to the inner part of the simulation box.

At the time shown, the magnetosphere is still forming and the filaments become longer. Close to the star the mass distribution is
quite complex, with an increased mass density at low latitudes close to the stellar surface and blobs of gas at low to mid-latitudes. Within the filament structure, gas moves back and forth along the field lines and inward towards the star. The mass distribution is distinctly asymmetric with respect to the equatorial plane. At this point we can not predict the final structure and extent of the magnetosphere but we notice the high mass density at low to mid-latitudes. To summarize, while there must be strong currents present to exert a Lorentz force strong enough to trap the gas, we do not find any changes of the field topology close to the star that would hint at possible reconnection events in that region.

5 DISCUSSION

Slow rotators with strong magnetic fields, such as NGC 1624-2, form dynamical magnetospheres, in which material flows along closed magnetic field loops from both magnetic poles, colliding near the magnetic equator and then falling back onto the stellar surface, leading to complex flows in the magnetosphere. NGC 1624-2’s magnetosphere was estimated to extend to an Alfvén radius of more than eleven stellar radii, hence trapping 95% of the outflowing wind, much more than other magnetic O stars that have an Alfvén radius of just a few stellar radii (David-Uraz et al. 2019). A much larger and denser magnetosphere compared to that of any other magnetic O-type star was also confirmed by Chandra observations, where the high X-ray luminosity, its variation with stellar rotation, and its large attenuation are consistent with a large dynamical magnetosphere with magnetically confined wind shocks (Petit et al. 2015).

NGC 1624-2 is the only magnetic O-type star for which the variability of both the longitudinal magnetic field and the mean magnetic field modulus is currently studied. The approximately sinusoidal variation of $\langle B_z \rangle$ and the ratio of the values of the extrema of $\langle B \rangle$ indicate that there is an important component of the field that is dipolar. The absence of a sign reversal of the mean longitudinal field over the rotation cycle indicates that only one magnetic pole is visible over the rotation cycle. Involving the mean field modulus measurements, we find that the geometrical structure of the magnetic field of NGC 1624-2 must depart from a centred dipole. A similar conclusion has been independently obtained by David-Uraz et al. (2020). Notably, among the eleven magnetic O-type stars with known $\langle B_z \rangle$ variation curves, for seven stars, HD 108, HD 37022 ($\omega^1$ Ori C), HD 54879, HD 148937, HD 191612, Tr 16-22, and NGC 1624-2, only one magnetic pole is visible throughout the rotation cycle (Shultz & Wade 2017; Petit et al. 2008; Hubrig et al. 2020; Wade et al. 2012b; Donati et al. 2006; Naze et al. 2016, this work). This implies that, due to an unfavorable inclination of the rotation axis, a larger fraction of their surfaces can never be observed, leaving the structure of the field over the invisible surface only constrained by the assumption of a tilted centred dipole. There will thus always be a considerable degree of ambiguity left in the magnetic models of these stars.

Our $\langle B_z \rangle$ values measured using He i emission lines are predominantly negative and, similar to the measurements of other spectral lines, show clear variability over the rotation period, with the negative $\langle B_z \rangle$ extremum observed around phase 0, where the photospheric absorption lines show a positive $\langle B_z \rangle$ extremum. To explain the presence of variable Stokes $V$ signatures corresponding to a longitudinal field of negative polarity in these lines, we suggest three alternative scenarios: one is related to a complex structure representing the contribution of several components, another one referring to the complex structure of the magnetospheric circumstellar environment of NGC 1624-2 with the possible presence of a reconnection-driven emission process, and a third one where the He i emission comes from circumstellar material accumulated in the equatorial region of the magnetosphere, where the dipolar magnetic field shows an inverted sign with respect to the pole.

As of today, magnetic reconnection is best studied in the solar photosphere, but it is not clear how the scenarios developed for the Sun can apply for massive stars. Although the detection of a negative longitudinal field inferred from the He i emission lines was reported about eight years ago, neither magnetohydrodynamical (MHD) simulations nor analytical models have been developed to provide a description of the magnetospheric structure permitting the reproduction of the observed phenomenon. According to Owocki et al. (2016) the presence of a huge magnetosphere and a very strong surface field are prohibitively expensive to model using numerical MHD. Our MHD numerical simulations show that while reconnection events indeed occur in magnetospheres of massive stars and the magnetic flux is carried away by the disk material, such events take place too far from the stellar surface where He i emission lines are formed.

On the other hand, as the shape of the Stokes $V$ profiles of the He i emission lines displayed in Figs. A4–A6 cannot be explained in terms of the Zeeman effect, the alternative hypothesis related to a composite structure of Stokes $I$ and $V$ profiles presenting the contribution of several components appears currently more promising. It is possible that future observations of all four Stokes parameters in these lines will help to disentangle the contribution of different line components.

ACKNOWLEDGEMENTS

We would like to thank the anonymous referee for their rigid review of our article, especially for their suggestions concerning the interpretation of the inverted magnetic field in the helium lines. We also thank G. Mathys for a discussion on the field modulus. SPJ is supported by the German Leibniz-Gemeinschaft, project number P67-2018. Based on observations collected at the Canada-France-

Figure 8. Mass density and magnetic field from a simulation with $B_0 = 1kG$. 

MNRAS 000, 1-9 (2020)
Hawaii Telescope (CFHT), which is operated by the National Research Council of Canada, the Institut National des Sciences de l’Univers of the Centre National de la Recherche Scientifique of France, and the University of Hawaii. Also based on archival observations obtained at the Bernard Lyot Telescope (Pic du Midi, France) of the Midi-Pyrénées Observatory, which is operated by the Institut National des Sciences de l’Univers of the Centre National de la Recherche Scientifique of France. Based on data acquired with the Potsdam Echelle Polarimetric and Spectroscopic Instrument (PEPSI) using the Large Binocular Telescope (LBT) in Arizona. This work has made use of the VALD database, operated at Uppsala University, the Institute of Astronomy RAS in Moscow, and the University of Vienna.

DATA AVAILABILITY
The ESPaDOnS data underlying this article are available in the CFHT Science Archive at https://www.cadc-ccda.hia-iha.nrc-cnrc.gc.ca/en/cfht/ and can be accessed with the object name. Similarly, the Narval data as well as the ESPaDOnS are available in PolarBase at http://polarbase.irap.omp.eu. The PEPSI and ARCES data underlying this article will be shared on a reasonable request to the corresponding author.

APPENDIX A: PROFILES FOR LONGITUDINAL MAGNETIC FIELD MEASUREMENTS
LSD Stokes $I, V$, and null profiles are presented in Fig. A1 for the line list with absorption lines, in Fig. A2 for the same line list, but without the Si absorption lines, in Fig. A3 for the high-excitation $\text{O} 
\text{i}\,\lambda7455$ emission line, in Fig. A4 for the He$\text{i}\,\lambda5876$ emission line, in Fig. A5 for the He$\text{i}\,\lambda6678$ emission line, and in Fig. A6 for the He$\text{i}\,\lambda7065$ emission line. The plots are sorted by rotational phase. Shaded regions in Stokes V and null panels indicate the mean uncertainty.

APPENDIX B: NEBULAR LINES
The strength of the forbidden [O$\text{i}$] nebular lines at $\lambda6300$ and $\lambda6363$ as a function of rotation phase is displayed in Fig. B1.

REFERENCES
Auer L. H., Mihalas D., 1972, ApJS, 24, 193
Caster J. I., Abbott D. C., Klein R. I., 1975, ApJ, 195, 157
Claret A., Bloemen S., 2011, A&A, 529, A75
Corcoran M., Ray T. P., 1997, A&A, 321, 189
David-Uraz A., et al., 2019, MNRAS, 483, 2814
David-Uraz A., Petit V., Shultz M. E., Fullerton A. W., Erba C., Keszthelyi Z., Seadrow S., Wade G. A., 2020, MNRAS, submitted (arXiv:2010.07482)
Donati J.-F., Semel M., Rees D. E., 1992, A&A, 265, 669
Donati J.-F., Semel M., Carter B. D., Rees D. E., Collier Cameron A., 1997, MNRAS, 291, 658
Donati J.-F., Babel J., Harries T. J., Howarth I. D., Petit P., Semel M., 2002, MNRAS, 333, 55
Donati J.-F., Howarth I. D., Bouret J.-C., Petit P., Catala C., Landstreet J., 2006, MNRAS, 365, 6
Finkenzeller U., 1985, A&A, 151, 340
Grunhut J. H., et al., 2009, MNRAS, L400, 94
Grunhut J. H., et al., 2017, MNRAS, 465, 2432
Howarth I. D., et al., 2007, MNRAS, 381, 433
Hubrig S., Nesvacil N., 2007, MNRAS, 378, L16
Hubrig S., Järvinen S. P., Schöller M., Hummel C. A., 2020, MNRAS, 491, 281
Kupka F., Dubernet M.-L., VAMDC Collaboration, 2011, Baltic Astronomy, 20, 503
Küker M., 2017, Astronomische Nachrichten, 338, 868
Kurtz D. W., Hubrig S., Järvinen S. P., Schöller M., 2020, Research Notes of the AAS, 4, 157
López Ariste A., 2002, ApJ, 564, 379
Mathys G., 1989, Fundam. Cosm. Phys., 13, 143
Muzerolle J., D’Alessio P., Calvet N., Hartmann L., 2004, ApJ, 617, 406
Nadezhdina I., Barbić J., Bagatto S., Morrell N., Gamen R., 2016, A&A, 596, A44
Owocki S. P., ud-Doula A., Sundqvist J. O., Petit V., Cohen D. H., 2016, MNRAS, 462, 3830
Petit P., Louge T., Théado S., Paletou F., Manset N., Morin J., Marsden S. C., Jeffers S. V., 2014, PASP, 126, 469
Petit V., Wade G. A., Drissen L., Montmerle T., Alecian E., 2008, MNRAS, 387, L23
Petit V., et al., 2013, MNRAS, 429, 398
Petit V., et al., 2015, MNRAS, 453, 3288
Preston G. W., 1967, ApJ, 150, 547
Ricker G. R., et al., 2014, Proceedings of the SPIE, 9143, 914320
Ryabchikova T., Piskunov N., Kurucz R. L., Stempels H. C., Heiter U., Pakhomov Y., Barklem P. S., 2015, Phys. Scr., 90, 054005
Sanchez Almeida J., Collados M., Del Toro Iniesta J. C., 1989, A&A, 222, 111
Schöller M., et al., 2017, A&A, 599, A66
Shultz M., Wade G. A., 2017, MNRAS, 468, 3985
Smith M. A., Cohen D. H., Hubeny I., Plettkamp S., Basri G., Johns-Krull C. M., MacFarlane J. J., Hirata R., 1997, ApJ, 481, 467
Stibbs D. W., 1950, MNRAS, 110, 395
Strassmeier K. G., et al., 2015, Astronomische Nachrichten, 336, 324
Strassmeier K. G., et al., 2018, Proceedings of the SPIE, 10702, 1070212
ud-Doula A., Owocki S. P., 2002, ApJ, 576, 413
ud-Doula A., Sundqvist J. O., Owocki S. P., Petit V., Townsend R. H. D., 2013, MNRAS, 428, 2723
Vink J. S., de Koter A., Lamers H. J. G. L. M., 2001, A&A, 369, 574
Wade G. A., et al., 2012a, MNRAS, 425, 1278
Wade G. A., et al., 2012b, MNRAS, 425, 2459
Walborn N. R., 1972, AJ, 77, 311
Wang et al., 2003, Proceedings of the SPIE, 4841, 1145
Ziegler U., 2008, Computer Physics Communications, 179, 227

This paper has been typeset from a TeX/Bibtex file prepared by the author.
Figure A1. LSD Stokes $I$, $V$, and null profiles for the line list with absorption lines. The plots are sorted by rotational phase. Shaded regions in Stokes $V$ and null panels indicate the mean uncertainty.
Figure A2. As Fig. A1, but for the line list without Si absorption lines.
Figure A3. Stokes $I$, $V$, and null profiles in velocity frame for the high-excitation O$_{III}$ λ7455 emission line (see Fig. A1).
Figure A4. As Fig. A3, but for the He II λ5876 emission line.
Figure A5. As Fig. A3, but for the He i λ6678 emission line.
Figure A6. As Fig. A3, but for the He\,\textsc{i} $\lambda$7065 emission line.
Figure B1. The strength of the forbidden [O i] nebular lines at $\lambda 6300$ (left) and $\lambda 6363$ (right) as a function of rotation phase.