On the relationship between the H$_2$ emission and the physical structure of planetary nebulae

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2012 June 30

ABSTRACT

Mid-IR observations of planetary nebulae (PNe) have revealed diffuse emission associated to their main nebular shells and outer envelopes or haloes. The interpretation of this emission is uncertain because the broad-band mid-IR images may include contributions of different components. In particular, the Spitzer IRAC 8 µm images, that best reveal these nebular features, can include contributions not only of H$_2$ lines, but also those of ionic species, PAH features, and thermal dust continuum emission.

To investigate the nature of the emission detected in mid-IR observations of a sample of 10 PNe, we have obtained narrow-band near-IR H$_2$ $\lambda$2.122 µm and optical [N ii] $\lambda$6584 Å images. The comparison between these images confirm that a significant fraction of the emission detected in the IRAC 8 µm images can be attributed to molecular hydrogen, thus confirming the utility of these mid-IR images to investigate the molecular component of PNe. We have also detected H$_2$ emission from PNe whose physical structure cannot be described as bipolar, but rather as ellipsoidal or barrel-like. These detections suggest that, as more sensitive observations of PNe in the H$_2$ $\lambda$2.122 line are acquired, the detection of H$_2$ emission is not exclusive of bipolar PNe, although objects with this morphology are still the brightest H$_2$ emitters. Finally, we remark that the bright H$_2$ emission from the equatorial ring of a bipolar PN does not arise from a photo-dissociation region shielded from the UV stellar radiation by the ring itself, but from dense knots and clumps embedded within the ionized material of the ring.

Key words: (ISM:) planetary nebulae: individual: A 66, M 1-79, M 2-48, M 2-51, NGC 650-51, NGC 6537, NGC 6563, NGC 6772, NGC 6778, NGC 7048 — I SM: jets and outflows — infrared: ISM — ISM: lines and bands

1 INTRODUCTION

Molecular hydrogen (H$_2$) can be expected in the photodissociation regions (PDR) of planetary nebulae (PNe), where the expanding envelope sweeps up the wind from the progenitor asymptotic giant branch (AGB) star. Molecular hydrogen can also form in neutral clumps embedded within the ionization zone, where high extinction, high density, and molecular (and dust) shielding allow H$_2$ to survive from the stellar radiation (McCandliss et al. 2007). The H$_2$ molecules in PNe can be excited by the ultraviolet (UV) radiation field of their central stars (CSPNe) in the photodissociation front (Black & van Dishoeck 1987) or by shocks (Burton, Hollenbach, & Tielens 1992). It has been suggested recently that H$_2$ can originally form in an excited state (Aleman & Gruenwald 2011).

It has been traditionally assumed that H$_2$ emission arises from regions where material is predominantly molecular. Therefore, it is not surprising that H$_2$ emission is mainly detected at the equatorial regions of bipolar PNe$^1$ because their thick equatorial disks would shield the UV radiation of the PN central star (CSPN), allowing the sur-

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$^1$ Hereafter we will adopt the definition of bipolar PNe given by Corradi & Schwarz (1995) as those whose Hα images display “an equatorial waist from which two faint, extended bipolar lobes depart”. When the morphology of the PN were insufficient to determine its physical structure (a ring-like PN can be interpreted as a pole-on bipolar source), we will rely only on kinematical information of the source.
vival of H$_2$ molecules (Kastner et al. 1996, Guerrero et al. 2004). Furthermore, the confinement of bipolar PNe to low Galactic latitudes have made them suspected to descend from the population of most massive progenitor stars of PNe (Corradi & Schwarz 1993). As these stars would eject thicker and more massive envelopes, this is considered an additional argument linking bipolar morphology with the significant presence of molecular material (Peimbert & Torres-Peimbert 1983, Hora & Latter 1996). The strong correlation between H$_2$ emission and bipolar morphology has originated the so-called Gatley’s rule (Kastner et al. 1994): “the detection of the 2.122 μm S(1) line of H$_2$ is sufficient to determine the bipolar nature of a PN.”

The increase in sensitivity of near- and mid-IR observations of H$_2$ lines (Hora 2006), and the access to other wavelength ranges (e.g., far-UV by FUSE, the Far-UV Spectroscopic Explorer) have revealed the presence of molecular hydrogen in PNe with a variety of morphologies (Dinerstein, Sterling, & Bowers 2006). The wavelength range between 1 and 10 μm is especially relevant because it includes a large number of transitions of molecular hydrogen (Turner & Zuckerman 1977). The intensity line ratios of some of these lines (the 2–1 S(1) λ2.2477/1–0 S(1) λ2.1218 and 1–0 S(0) λ2.2235/1–0 S(0) λ2.1218 particularly) can be used to infer the molecular excitation mechanism (shocks or UV radiation field) and the physical conditions of H$_2$ (see Likkel et al. 2006).

The four Spitzer Infrared Array Camera (IRAC) bands include H$_2$ lines such as the 1–0 (0(5) λ3.235 μm and 0–0 S(13) λ3.8464 μm lines in the IRAC 3.6 μm band, the 0–0 S(1) λ4.181 μm and 0–0 S(9) λ4.6947 μm lines in the IRAC 4.5 μm band, the 0–0 S(7) λ5.5115 μm and 0–0 S(6) λ6.1088 μm lines in the IRAC 5.8 μm band, and the 0–0 S(5) λ6.9091 μm and 0–0 S(4) λ8.0258 μm lines in the IRAC 8 μm band. Observations of PNe have shown the relevance of nebular emission in the IRAC 8 μm band as it contains few emission from stars and more intense diffuse emission than the IRAC 3.6 and 4.5 μm bands (e.g., Anderson et al. 2012), while having a better sensitivity than the IRAC 5.8 μm band. Furthermore, the IRAC 3.6 and 4.5 μm bands are dominated by bremsstrahlung emission and by the ionic transition lines of Br$\alpha$ λ4.052 μm, [Mg iv] λ4.49 μm, and [Ar vi] λ4.53 μm (Phillips & Ramos-Larios 2010). IRAC 8 μm images of PNe frequently reveal emission from extended haloes and molecular knots (Ramos-Larios & Phillips 2009, Chi 2011, Phillips & Marquez-Lugo 2011). The comparison between IRAC 8 μm and H$_2$ images of PNe (e.g., NGC 6720 and NGC 7293) reveals an excellent match between the appearance in the two bands that indicates a common origin (Hora 2006, Hora et al. 2006), although molecular hydrogen is suspected, the contribution from the [Ar iii] λ8.991 μm line cannot be neglected (Hora et al. 2004).

Accordingly we have searched the Spitzer archive for IRAC 8 μm images of PNe showing extended emission that could be attributed to molecular material. The sample, composed of 8 morphologically diverse PNe, has subsequently been imaged through narrow-band optical and near-IR filters to ascertain the presence and spatial distribution of ionized material and H$_2$. Two PNe, namely M1-79 and NGC 6778, were added to this sample as they exhibit intriguing narrow-band optical and near-IR morphologies, despite there are no Spitzer images for them. The comparison between optical, near- and mid-IR images has allowed us to verify the presence of molecular material in regions of H$_2$ emission, but also from regions where H$_2$ is shielded from the stellar UV radiation field, so that the H$_2$ molecules are neither disrupted nor excited. In this paper we present evidence that the bipolar PNe–H$_2$ relationship is not as close as claimed by Gatley’s rule, as H$_2$ emission is found in PNe with morphological types other than bipolar.

The observations and archival data are presented in Section 2. The results for each individual PN are described in Section 3. The discussion and final summary are presented in Section 4.

## 2 OBSERVATIONS AND ARCHIVAL DATA

### 2.1 Optical images

Most of the optical images presented in this paper (Table I) have been obtained at the Observatorio Astronómico Na-
cional (OAN, Mexico), using either the 1.5m Harold Johnson or the 0.84m telescopes. At the 1.5m telescope, a [N ii] filter (λ = 6584 Å, Δλ = 11 Å) was used, whereas at the 0.84m telescope either a [N ii]+Hα filter (λ = 6564 Å, Δλ = 72 Å) or a [N ii] filter (λ = 6583 Å, Δλ = 10 Å) were used. The narrow-band [N ii] images of M 1-79, M 2-48, M 2-51, and NGC 7048 were obtained using the RUCA filter wheel (Zazueta et al. 2010). The detector was a 1024 × 1024 CCD that provides a pixel scale of 0.′39 pixel\(^{-1}\), binning 1×1 and a FoV of 6′.7 × 6′.7. The [N ii]+Hα image of NGC 650-51 was obtained using SOPHIA (Sistema Óptico Para Hacer Imágenes de campo Amplio), an optical system for the acquisition of wide-field images at the 0.84m OAN telescope. This time, the detector was a 2048×4608 CCD e2v, named ESOPO, providing a pixel scale of 0′.07 pixel\(^{-1}\), binning 1×1 and a FoV of 30′0 × 30′0.

For NGC 6778, we have used the [N ii] images published by Miranda, Ramos-Larios, & Guerrero (2010). The images were obtained through the narrow-band [N ii] filter (λ = 6584 Å, Δλ = 9 Å) using ALFOSC (Andalucía Faint Object Spectrograph and Camera) in its imaging mode at the 2.56m Nordic Optical Telescope (NOT) of the Observatorio del Roque de los Muchachos (ORM), La Palma, Spain. The e2v 2048×2048 CCD detector used for these observations provides a pixel scale of 0′.184 pixel\(^{-1}\) and a FoV of 6′.3 × 6′.3.

Finally, for A 66 we used an R band POSS-II image downloaded from the ESO Digitized Sky Survey (DSS) with pixel scale ≃ 0′.0 pixel\(^{-1}\).

2.2 Near-IR images

The narrow-band H₂ λ2.122 μm and K\(_s\) continuum near-IR images were mainly obtained at the 4.2m William Herschel Telescope (WHT) of the ORM using LIRIS (Long-Slit Intermediate Resolution Infrared Spectrograph, Acosta Pulido et al. 2003). The detector was a 1024×1024 HAWAII array with plate scale 0′.25 pixel\(^{-1}\) and FoV of 4′.3 × 4′.3. The narrow-band H₂ λ2.122 μm and K\(_s\) continuum λ2.27 μm images of NGC 7048 were obtained using NICS (Near Infrared Camera and Spectrometer, Oliva & Gennari 1997) at the 3.5m Telescopio Nazionale Galileo (TNG) of the ORM. The Rockwell Hawaii 1024×1024 array used for these observations has a projected scale of 0′.25 pixel\(^{-1}\) and a FoV of 4′.3 × 4′.3. The narrow-band H₂ λ2.122 μm image of A 66 was obtained using the SOFI camera (Son OF Isaac, Moorwood, Cuby, & Lidman 1998) at the 3.5m New Technology Telescope (NTT) of La Silla Observatory, Chile. The detector was a 1024×1024 HgCdTe HAWAII array providing in its large field mode (LF) a pixel scale of 0′.288 pixel\(^{-1}\) and a FoV of 4′.9 × 4′.9. Finally, Two Micron All Sky Survey (2MASS) JHK\(_s\) images with pixel scale 1′.0 pixel\(^{-1}\) are presented for M 2-51.

The DSS is an all-sky photographic survey conducted with the Palomar and UK Schmidt telescopes. The Catalogs and Surveys Branch (CASB) is digitizing the 0′.5 × 0′.5 photographic plates using a modified PDS microdensitometer to support HST observing programs and to provide a service to the astronomical community.
2.3 Mid-IR images

The mid-IR images used in this paper have been downloaded mostly from the Spitzer archives. The Spitzer IRAC images of M2-48 and NGC 6537 belong to the Galactic Legacy Infrared Mid-Plane Survey Extraordinaire (GLIMPSE) program which used IRAC to map the Galactic plane in the range $|l| \leq 60^\circ, |b| \leq 1^\circ$ (Fazio et al. 2004). The GLIMPSE images in the IRAC 8 $\mu$m band used in this paper have a spatial resolution $\simeq 2^\prime$. Similarly, we have used the Spitzer IRAC 8 $\mu$m images of A66 and NGC 7048 (Program ID 30285, Spitzer Observations of Planetary Nebulae 2, PI: Giovanni Fazio), NGC 650-51, NGC 6563, and NGC 6772 (Program ID 68, Studying Stellar Ejecta on the Large Scale using SIRTF-IRAC, PI: Giovanni Fazio), and M2-51 (Program ID 50398, Spitzer Mapping of the Outer Galaxy, SMOG, PI Sean Carey). The spatial resolution of these images varies between $\simeq 1.7$ and $\simeq 2.0$.

No Spitzer images are available for M1-79 and NGC 6778. For these nebulae, we have used WISE (Wide-field Infrared Survey Explorer, Wright et al. 2010) images retrieved from the NASA/IPAC Infrared Science Archive (IRSA). WISE is a NASA Explorer mission to surveys the entire sky at 3.4, 4.6, 12, and 22 $\mu$m, the so-called W1 through W4 bands, with 5$\sigma$ point source sensitivities better than 0.08, 0.11, 1, and 6 mJy, respectively. The 40cm telescope uses HgCdTe and Si:As detectors arrays with a plate scale of 2.75 pixel$^{-1}$. The W2 4.6 $\mu$m images were downloaded from the WISE All-Sky Data Release. The images have angular resolution $\simeq 6.4''$ and astrometric accuracy for bright sources better than 0.15.

2.4 Spitzer Spectroscopy in the MIR

Spectroscopic observation used in this paper were acquired using the Short-Low (SL) module 1 (SL1) and module 2 (SL2) at short (5.1-8.5 $\mu$m) and long (7.4-14.2 $\mu$m) wavelength respectively of the Spitzer Infrared Spectrograph (IRS; Houck et al. 2004). The spectra of M2-51 and NGC 6537 were obtained through Spitzer Program 45 (Deuterium Enrichment in PAHs; P.I. Thomas Roellig) on 1/6/2004 and Spitzer Program 50179 (Planetary Nebulae As A Laboratory For Molecular Hydrogen in the Early Universe; P.I. Kris Sellgren) on 4/11/2008 respectively.

3 RESULTS

Based on the inspection of the optical images and previous spatio-kinematical studies, when available, we have divided our sample into two broad morphological groups: elliptical PNe and bipolar PNe. The first group includes A66, M2-51, NGC 6563, NGC 6772, and NGC 7048, whereas the group of bipolar PNe is composed of M1-79, M2-48, NGC 650-51, NGC 6537, and NGC 6778.

3.1 Elliptical PNe

3.1.1 A66 — PN G019.8−23.7

A66 was included in the list of old, evolved PNe compiled by Abell (1955, 1966). The DSS optical image (blue in Figure 1 left) shows a roughly spherical, low surface brightness shell of radius $\simeq 118''$. The best quality narrow-band optical images of A66 were presented by Hua, Dopita, & Martins (1998) who described it as a roundish, old PN with a blowout structure towards the NE. They also reported the presence of radial structures or filaments escaping outwards, and a band of emission in H$\alpha$ and [N II] crossing the central regions and dividing the nebula into two cavities.

The morphology of A66 in the H$\alpha$ and IRAC 8 $\mu$m images highlights the radial structures hinted in optical images. The H$\alpha$ image (Figure 1 right) shows a series of cometary knots that are mostly distributed along a central band and
a fragmented ring of radius 133″ broken towards the northeast. There is a clear correlation between these morphological features and those described in optical images: the central band is coincident with that observed in the optical, the ring-like structure encompasses the optical emission, and the lack of H₂ emission towards the northeast corresponds with the Hα and [N ii] blowout.

The IRAC 8 μm image (red in Figure 1 left) shows very similar morphology in the central regions to that of the H₂ image, but its larger FoV reveals emission extending farther out. Indeed, the H₂ cometary knots that overrun the NTT H₂ image stretch out in the IRAC 8 μm emission up to radial distances ≃240′′. The central optical nebula is surrounded by a halo of emission in the IRAC 8 μm image. The similarities between H₂ and IRAC 8 μm emission in the central regions and the identification of some of the outermost features in the IRAC 8 μm image with radial knots in the H₂ image strongly suggests that the outermost emission detected in the IRAC 8 μm image is produced by H₂ molecules.

3.1.2 M 2-51 — PNG103.2+00.6

Optical [N ii] images of M 2-51 (Jewitt, Danielson, & Kupferman 1986; Balick 1987) have revealed an elliptical morphology with a size ≃36″×56″, and major axis almost along the north-south direction. Our [N ii] image detects this elliptical shell, as well as an external elliptically shaped outer shell of size ≃60″×86″ whose major axis is tilted by ≃30° with respect to that of the inner shell. Low surface brightness diffuse emission is also detected along the minor axis of the outer elliptical shell up to a radial distance of 40″.

No narrow-band observations of the H₂ λ2.122 μm line are available in the literature for M 2-51. We have compared the 2MASS Ks image with those in the J and H bands to search for a photometric excess that could be used as a proxy for detection of H₂ emission (see Ramos-Larios, Kemp, & Phillips 2006, for details on this technique). The 2MASS image (Figure 2 right), similar in quality to the JHK images presented by Saito et al. (1999), reveals hints of emission excess from a filamentary structure consistent in size and location with the outer elliptical shell, but the low signal of this emission does not allow us to make a firm statement.

The image in the IRAC 8 μm band (red in Figure 2 left) generally follows the double shell morphology and diffuse emission hinted in the [N ii] image. The outer elliptical shell has a size ≃80″×116″. The ionic or molecular nature of this emission is uncertain.

3.1.3 NGC 6563 — PNG358.5−07.3

NGC 6563 was included in the catalogue of narrow-band images of PNe of Schwarz, Corradi, & Melnick (1992). Its Hα+[N ii] image displays a main body with elliptical shape and two extensions or ansae. Our [N ii] image (blue in Figure 3 left) shows similar elliptical morphology oriented along the northeast-southwest direction (major axis along P.A. ≃50°) with a size of 38″×52″. The two ansae protrude from the bright inner shell almost along the east-west direction up to radial distances ≃40″. Whereas the optical morphology of NGC 6563 may be interpreted as a wide equatorial belt and narrow bipolar lobes seen almost pole-on, the kinematics do not to confirm this interpretation (Vázquez et al., in prep.), but rather confirm an ellipsoidal structure with short, low velocity extensions consistent with ansae (Stanghellini, Corradi, & Schwarz 1992).

The distribution of molecular hydrogen in NGC 6563
is revealed for the first time in our H$\alpha$ λ2.122 µm image (green in Figure 3-left and Figure 3-right). The molecular emission outlines that of the elliptical ionized region, both showing a pattern of spiral-like dark lanes and bright filaments. These features are typical of bright ring-like PNe such as the Ring and the Helix Nebula and have been suggested to form part of a tilted barrel-like structure (Meixner et al. 2003; O’Dell, McCullough, & Meixner 2004; Speck et al. 2002, 2003). The molecular and ionized emissions in the ansae, however, differ notably: in [N ii], the emission is uniform, fills the ansae, and falls off with radial distance, whereas the H$_2$ emission encloses the [N ii] emission, delineating the ansae edges with a remarkable point-symmetric brightness distribution. No H$_2$ emission is detected outside the bright inner shell along the north-south direction, where the walls of this shell are thicker and may imply more efficient shielding from the stellar UV radiation. On the other hand, weak, diffuse emission is detected forming a broken, round shell of radius $\approx 50''$. The emission is notably limb-brightened along an arc towards the east and northeast, but it appears fuzzy towards the opposite side of the main nebula.

The emission from the bright optical and H$_2$ shell and ansae of NGC 6563 are detected in the IRAC 8 µm image (Figure 4-left), revealing more clearly the outer shell of size 105'' that surrounds the inner elliptical shell and its ansae. This outer shell is well defined along the northeast half, but its appearance is more diffuse in its southwest half. Otherwise, the morphology in the IRAC 8 µm image of the inner shell is similar to that of the ionized and hydrogen molecular material, but there are some subtle differences: the IRAC 8 µm image highlights a pattern of filaments inside the elliptical shell and the ansae are both broader and extend further. Whereas the origin of the IRAC 8 µm emission of the inner shell can be attributed to ionized and (most likely) H$_2$ molecular lines, the nature of the material in the outer shell appears to be H$_2$ emission, although some contribution of thermal dust emission cannot be excluded (e.g., Phillips et al. 2009).

### 3.1.4 NGC 6772 — PN G033.1−06.3

NGC 6772 appears in a large number of optical imaging studies of PNe (Jewitt, Danielson, & Kupferman 1986; Jacoby & Kaler 1989; Schwarz, Corradi, & Melnick 1992; Bachiller et al. 1993; Zhang & Kwok 1998). Our [N ii] image (blue in Figure 4-left) confirms the barrel-like elliptical morphology previously described (e.g., Jewitt, Danielson, & Kupferman 1986). The outer edge of this thick elliptical shell, oriented mostly along the north-south direction, has a size of 31'' x 44''. The shell is distorted along the northeast-southwest direction, where two ansae or blisters protrude from the shell. Our [N ii] image confirms the presence of outer emission, mostly distributed along the east-west direction, but it also reveals a new structure, an arc of radius $\approx 66''$ towards the east that appears to be one half of an outer round shell.

Previous studies of the spatial distribution of H$_2$ in NGC 6772 (Webster et al. 1988) showed an elliptical shell distorted towards the northeast and southwest regions. Our H$_2$ λ2.122 image (Figure 4-right) reveals a wealth of details in this elliptical shell, as well as series of features outside it. The radial features protruding from the inner shell are certainly remarkable. The arc-like feature hinted in the [N ii] images towards the East of the main nebular shell is clearly detected in H$_2$ with a similar radius, $\approx 66''$. This structure has a notable limb-brightness morphology towards the east, whereas it fades and extends further out towards the west. Overall, this morphology is reminiscent of a shell interacting with the Interstellar Medium (ISM), either by the nebular proper motion or by density gradients in the ISM (Ramos-Larios & Phillips 2004). Similar morphology can be
claimed for NGC 6563, but the case of NGC 6772 is certainly more clear.

The IRAC 8 μm emission (red in Figure 4) shows the same spatial distribution than the H$_2$ emission in the inner elliptical shell, including the distorted regions towards the northeast and southwest. The radial features described in H$_2$ are also detected in the 8 μm image, but while the H$_2$ rays are concentrated just outside the inner elliptical shell, in the IRAC 8 μm image the rays are more evenly distributed around the ellipse and extend at greater distances. This coincidence can be interpreted as a common origin for this emission i.e., molecular hydrogen, whereas the outer section of the rays not detected in the H$_2$ image may imply that H$_2$ is present but shielded from UV radiation, so that it is not excited to emit significantly in the 1–0 S(1) A$_2.122$ line. The outer shell is more clearly revealed in this IRAC 8 μm image than in the H$_2$ band, with a notable bow-shock morphology towards the northeast. Along the opposite direction, the IRAC 8 μm emission is diffuse and suggests it is trailing the main nebula that would be moving with respect to the ISM.

3.1.5 NGC 7048 — PN G088.7–01.6

Despite being an extended, relatively bright PN, the morphology and physical structure of NGC 7048 have not been studied in detail. The most recent narrow-band optical images mapping its ionized component were presented by Balick (1987) who described it as a middle elliptical PN, a conclusion also reached in other studies based on the same sets of images (Zhang & Kwok 1998). Our [N II] image (blue in Figure 5) shows a filamentary, almost round shell of radius $\approx 30''$ marked by bright eastern and western arcs that leave an opening towards the north and south. Extended emission is detected along these directions at longer radial distances, up to $\approx 40''$ towards the north and $\approx 35''$ towards the south. Rather than an elliptical shell, NGC 7048 resembles a tilted barrel that opens at the poles (Frank et al. 1993). Our [N II] image reveals a weak limb-brightened round shell $\approx 55''$ in radius that can be described as a halo. This halo is not completely concentric with the bright inner shell, neither its surface brightness is azimuthally constant: the halo has two bright arcs, north and south of the inner shell, along the directions of their polar openings.

The emission in the H$_2$ $\lambda 2.122$ μm line has been previously described by Kastner et al. (1996) and Davis et al. (2003). Both works report bright H$_2$ emission from a filamentary barrel-like structure, but the larger FoV images of Kastner et al. (1996) unveil emission along P.A. 10$^\circ$, i.e., mostly along the north-south direction. Indeed, this emission is similar to that shown in our [N II] image. A close comparison with our H$_2$ image (Figure 5) confirms these similarities, but it also shows evidences that the diffuse emission emanating through the north and south openings of the inner shell is relatively brighter in the [N II] image (blue color in Figure 5). As the spectroscopic study of Davis et al. (2003) for the excitation mechanism of the H$_2$ emission suggests, the emission from the main shell is shock-excited, probably from the propagation of a low velocity shock generated by the inner shell expanding into the outer halo (Medina et al. 2007).

The IRAC 8 μm image (red color in Figure 5) follows the H$_2$ and [N II] filaments and extended emission of the inner shell of NGC 7048, but where these emissions are faintly detected in the outermost regions, the emission in the IRAC 8 μm band is intense and clearly reveals a round shell of radius $\approx 56''$. The emission in this band shows the limb-brightened morphology, but it also discloses radial bright and dark stripes and an azimuthally dependent brightening. These can be associated to the shadowing of the central star by the eastern and western arcs of the inner shell that are most likely optically thick to the UV radiation from the
central star, probably in the cooling track of white dwarfs. The correspondence between the H$_2$ and [N II] features and those in the IRAC 8 µm image suggests that the emission in the inner shell and halo detected in this latter band includes significant contributions of emission lines both from ionized material and molecular hydrogen.

3.2 Bipolar PNe

3.2.1 M 1-79 — PN G093.3—02.4

Narrow-band images of M 1-79 were presented by Manchado et al. (1996) and its morphology, kinematics, and physical structure in optical emission lines of the ionized gas has been extensively studied by Saurer (1997). This latter work describes M 1-79 as a 46″ ×24″ bright elliptical shell with its major axis oriented near the east-west direction (P.A.≃85°). A bright bar crosses the shell at P.A.≃14°, i.e., this bar is misaligned with respect to the ellipse minor axis and reminds the so-called ‘big-tail’ at the central region of the bipolar PN NGC 2818 (Vázquez et al. 2012). A pair of claw-like features protrude from the bright central region along the southeast-northwest (P.A.≃140°) direction. A high contrast image reveals bipolar lobes that extend further out, up to distances ≃45″.

Our [N II] image (Figure 3-left) confirms these structural components, additionally revealing that the outer pair of bipolar lobes are tilted with respect to the claw-like features and that there is an even larger northwestern bipolar lobe which extends up to ≃65″. We also detect faint diffuse emission towards the east of the bright inner shell, but there is no clear evidence it takes part of a complete outer shell.

The narrow-band H$_2$ λ 2.122 image of M 1-79 (green in Figure 3-left and Figure 5-right) discloses for the first time the molecular hydrogen distribution in this PN which is in many aspects different from the distribution of ionized material. The outskirts of the optically bright inner shell are delineated in H$_2$, but there are no signs of the bipolar lobes. On the contrary, we detect in H$_2$ a series of bright radial filaments and shadows emanating from the bright inner shell that, avoiding the directions along which the bipolar lobes are detected, are enclosed within an ellipse of size ≃42″ ×55″. We note that some of the brightest filaments in H$_2$ are spatially coincident with the diffuse [N II] emission detected towards the east of the bright inner shell. We also note that the bright bar in the inner shell produces a remarkable conical shadow in the H$_2$ emission, thus suggesting that there is not enough UV flux to excite the H$_2$ molecules along these directions.

Unfortunately, there are no available Spitzer IRAC images of M 1-79. We have thus used the WISE W2 4.6 µm image to investigate the properties of this nebula at longer wavelengths. This image (red in Figure 8-right) shows emission from the central regions of M 1-79, but the limited spatial resolution and sensitivity of WISE, and the possible contribution of near-IR ionic lines to the W2 band are not adequate to study the molecular component of the outer regions of this nebula in this mid-IR image.

3.2.2 M 2-48 — PN G062.4—00.2

Different morphological (Corradi & Schwarz 1995; Manchado et al. 1996) and kinematical (Vázquez et al. 2003; López-Martín et al. 2002; Dobrinič et al. 2008) studies of M 2-48 (a.k.a. Hen 2-449) have revealed a highly collimated bipolar PN with an obscured waist, bow-shock features along its major axis, and a fragmented off-center round shell. Our [N II] image (blue in Figure 4-left) confirms the bow tie-shaped core with size ≃9″ ×15″ and detects a bow-shock feature east of the main nebula along the major axis at ≃55″, and two outer bow-shock features ≃95″ east and ≃120″ west along an axis tilted by +5° with respect to the bipolar axis of the main nebula. The eastern bow-shock feature at 55″ is coincident with the fragmented round shell that shows a radius of 45″ and is off-centered by 14″ towards the northeast.

Our H$_2$ λ 2.122 image (Figure 7-right) shows faint emission encompassing the ionized bipolar lobes, but no H$_2$ emission is detected at the nebular core. There is also even fainter H$_2$ emission just interior of the brightest eastern and western arcs of the fragmented off-center round shell, and the eastern bow-shock feature at 55″. The latter is suggestive of shock excitation of the H$_2$ molecules at these locations.

The IRAC 8 µm image of M 2-48 has been described by Kwok et al. (2008) and Phillips & Ramos-Larios (2008). Here we note that this mid-IR image shows bright emission at the core of the main nebular shell and faint emission associated to the bipolar lobes. The emission from the bipolar lobes in this image has a biconical morphology and follows more closely the H$_2$ emission than the [N II] emission, thus indicating that it corresponds to emission from H$_2$ molecules. No emission at 8 µm seems to be associated to the outer round shell nor to the bow-shock features, but we note that M 2-48 is embedded within a region of patchy, diffuse emission.

3.2.3 NGC650-51 — PN G130.9—10.5

The optical and mid-IR properties of NGC650-51 (a.k.a. M 76) have been recently studied in detail by Ramos-Larios, Phillips & Cuesta (2008, see also references therein). Our [N II] image (blue in Figure 8-left) shows a bright, tilted ring-like structure oriented along the northeast to southwest direction with an angular size of 66″ ×144″. Two bipolar lobes with archetypical butterfly morphology extend along the minor axis of the central ring up to distances ≃85″. At the tips of these lobes, fainter bow-shaped structures are detected, increasing the total extent of the bipolar lobes up to 260″.

Previous observations of NGC650-51 in the H$_2$ line emphasized the presence of diffuse emission from the central ring, which seemed to be brighter at the tips of its long axis (Kastner et al. 1996). Our H$_2$ image (Figure 8-right) resolves this ring into a series of disconnected knots and filaments. Some knots and filaments are also detected in the inner regions of the bipolar lobes.

The spatial distribution of the emission of NGC650-51

\[ \text{K}_c \] The \text{K}_c image used to obtain the continuum free H$_2$ image in Figure 4-right also shows bright emission at the nebular core.
Relationship between $H_2$ and structure of PNe

3.2.4 NGC 6537 — PN G010.1+00.7

NGC 6537 is a high velocity bipolar PN with a noticeable point-symmetric brightness distribution of the bipolar lobes. In the IRAC 8 μm image (red in Figure 8-left) has been compared to those of the ionized and $H_2$ material (see also Hora et al. 2004; Ramos-Larios, Phillips, & Cuesta 2008). The emission in this mid-IR band closely follows the [N ii] image, showing the central ring, the inner bipolar lobes, and their fainter extensions. In addition, the mid-IR emission highlights faint filaments outside the main nebular body.

in the IRAC 8 μm image (red in Figure 8-left) has been compared to those of the ionized and $H_2$ material (see also Hora et al. 2004; Ramos-Larios, Phillips, & Cuesta 2008). The emission in this mid-IR band closely follows the [N ii] image, showing the central ring, the inner bipolar lobes, and their fainter extensions. In addition, the mid-IR emission highlights faint filaments outside the main nebular body.

in the IRAC 8 μm image (red in Figure 8-left) has been compared to those of the ionized and $H_2$ material (see also Hora et al. 2004; Ramos-Larios, Phillips, & Cuesta 2008). The emission in this mid-IR band closely follows the [N ii] image, showing the central ring, the inner bipolar lobes, and their fainter extensions. In addition, the mid-IR emission highlights faint filaments outside the main nebular body.

in the IRAC 8 μm image (red in Figure 8-left) has been compared to those of the ionized and $H_2$ material (see also Hora et al. 2004; Ramos-Larios, Phillips, & Cuesta 2008). The emission in this mid-IR band closely follows the [N ii] image, showing the central ring, the inner bipolar lobes, and their fainter extensions. In addition, the mid-IR emission highlights faint filaments outside the main nebular body.

Figure 6. WISE W2 4.6 μm (red), WHT $H_2$ λ2.122 (green), and OAN [N ii] (blue) colour-composite RGB picture (left), and WHT continuum-subtracted $H_2$ λ2.122 image (right) of M 1-79.

Figure 7. IRAC 8 μm (red), WHT $H_2$ λ2.122 (green), and OAN [N ii] (blue) colour-composite picture (left), and WHT continuum-subtracted $H_2$ λ2.122 image (right) of M 2-48.

Figure 8. WISE W2 4.6 μm (red), WHT $H_2$ λ2.122 (green), and OAN [N ii] (blue) colour-composite RGB picture (left), and WHT continuum-subtracted $H_2$ λ2.122 image (right) of M 1-79.

Figure 9. IRAC 8 μm (red), WHT $H_2$ λ2.122 (green), and OAN [N ii] (blue) colour-composite picture (left), and WHT continuum-subtracted $H_2$ λ2.122 image (right) of M 2-48.

Figure 10. WISE W2 4.6 μm (red), WHT $H_2$ λ2.122 (green), and OAN [N ii] (blue) colour-composite RGB picture (left), and WHT continuum-subtracted $H_2$ λ2.122 image (right) of M 1-79.
brightened edges of the bipolar lobes. The detection of H$_2$ emission at the nebular core is uncertain, as $K_s$ images also show bright emission at this location, but the images presented by Davis et al. (2003) seem to confirm a ring-like feature of H$_2$ emission. Our continuum-subtracted H$_2$ image (Figure 9-right) displays this ring, but it also traces the faintest emission from the bipolar lobes. Relatively strong H$_2$ emission is detected at the tip of the northeastern bipolar lobe which is otherwise rather faint in the [N ii] image.

The IRAC 8 µm image of NGC 6537 (red in Figure 9-left, see also Kwok et al. 2008; Phillips & Ramos-Larios 2008) shows a bright unresolved source at the nebular core and faint diffuse emission that traces the inner regions of the bipolar lobes. At least for the northeastern lobe, its tip is detected in the 8 µm band.

3.2.5 NGC 6778 — PN G034.5–06.7

NGC 6778 had received little attention until the discovery of a binary CSPN (Miszalski et al. 2011) and a disrupted equatorial ring fragmented by fast stellar winds and multiple collimated outflows (Guerrero & Miranda 2012). The [N ii] image (blue in Figure 10-left, adopted from Guerrero & Miranda 2012) has been compared to the H$_2$ λ2.122 µm (green in Figure 10-left) and continuum-subtracted H$_2$ (Figure 10-right) images. The H$_2$ emission traces the brightest [N ii] emission at the tips of the equatorial regions. This spatial distribution is reminiscent of a barrel-like structure and may imply that the object is density-bounded along the equatorial plane.

Unfortunately, there are no available Spitzer images of NGC 6778. The WISE W2 4.6 µm image (red in Figure 10-left) shows a bright, unresolved source at the location of the central regions of NGC 6778. The limited spatial resolution and sensitivity of WISE, and the possible contribution of near-IR ionic lines to the W2 band are not adequate to study the molecular component of the outer regions of this nebula.

4 DISCUSSION

4.1 Interpreting the emission in IRAC 8 µm images

Multiple studies have shown the efficiency of IRAC images in the 8 µm band (6.2994 µm ≤ λ ≤ 9.5876 µm) to detect extended haloes around PNe (e.g., Ramos-Larios & Phillips 2009) and dense knots embedded within ionized nebular shells (e.g., Hora et al. 2006). This emission can be attributed to the contribution into the IRAC 8 µm bandpass of the H$_2$ 0–0 S(5) λ6.0991 µm, 1–1 S(5) λ7.297 µm, and 0–0 S(4) λ8.0258 µm rotational lines. However, the contribution of other emission lines from ionic species, such as [Ar ii] λ6.985 µm, [Ne iv] λ7.642 µm, [Ar v] λ7.902 µm, and [Ar iii] λ8.991 µm, cannot be neglected for the inner ionized nebular regions (Hora et al. 2004). Similarly, the contribution of the PAH features at 6.2, 7.7 and 8.6 µm may be of importance for dusty regions, such as obscured equatorial waists of bipolar PNe and dense knots.

The comparison between H$_2$ and IRAC 8 µm images of PNe in our sample reveals a close correlation between morphological features seen in the two bands for a significant fraction of sources. The central regions and outermost shells of A 66, NGC 6563, NGC 6772, and NGC 7048 show overall extended haloes around PNe and dense knots. The H$_2$ and IRAC 8 µm images of A 66 display a fragmentary ring and a series of small-scale structures with an appearance of cometary knots. These knots, whose heads are clearly detected in the innermost regions of A 66 mapped by the H$_2$ image, extend further out in long filaments that are very notable in the larger FoV of the IRAC 8 µm image. Similarly, the inner nebular regions of A 66 mapped by the H$_2$ image are not adequate to study the molecular component of the outer regions of this nebula.
lae of NGC 6563, NGC 6772, and NGC 7048 show an excellent correspondence in the H$_2$ and 8 µm images, even on small-scale filaments and knots. The H$_2$ images of these PNe are indicative of limb-brightened, faint outer shells that are brighter and revealed as complete envelopes in the IRAC 8 µm images. These shells, which can be described as haloes (Chu, Jacoby, & Arendt 1987), are spatially coincident in the H$_2$ and 8 µm images. As for A66, there is a series of bright (rays) and dark (shadows) radial filaments that connect the inner and outer shells of NGC 6772.

The similar spatial distribution of the H$_2$ and IRAC 8 µm images in these PNe suggests that a significant fraction (if not all) of the emission in the IRAC 8 µm images of these PNe can be attributed to lines of molecular hydrogen in this IRAC filter bandpass. This conclusion is supported by the Spitzer IRS spectra of NGC 6720 and NGC 7293 which show that the IRAC 8 µm is dominated by H$_2$ emission lines (Hora et al. 2003, 2006, 2009). To test this suggestion, we have examined the Spitzer IRS spectra in the 5.8-8.0 µm band available for the PNe in our sample, namely M 2-51 and NGC 6537. The spectra presented in Figure 11 shows that the mid-IR emission in the IRAC 8 µm band from M 2-51 and the bipolar lobes of NGC 6537 present prominent H$_2$ emission of the transitions 0–0 S(5) λ6.9991 µm, 1–1 S(5) λ7.2801 µm, and 0–0 λ8.2058 µm. On the other hand, the mid-IR emission from the central regions of NGC 6537 is dominated by emission lines of ionic species such as [Ne iv] λ7.642 µm and [Ar iii] λ8.991 µm, and the PAH feature at 7.7 µm.

We are thus confident that the extended, outermost emission detected in the IRAC 8 µm images can be attributed to H$_2$, whereas some contribution of emission lines from ionized material can be expected in the innermost regions. The H$_2$ emission in the PNe in our sample is mostly associated to shell-like structures and its excitation may be two-fold as discussed below. The inner shell of NGC 7048 is shock-excited (Davis et al. 2003), probably associated to the propagation of a small-velocity shock generated by the expansion of the inner shell into the outer shell. The emission of H$_2$ in the inner shells of NGC 6563 and NGC 6772 may be similarly shock-excited, whereas the H$_2$ emission from the outer shells of A66, NGC 6772 and NGC 7048 seems to exhibit a dependence with “openings” in the inner shell that is suggestive of UV excitation.

Contrary to these nebulae, the spatial distributions of the H$_2$ and IRAC 8 µm emissions in the bipolar PNe M 2-48, NGC 650-51, and NGC 6537, do not correlate closely. The H$_2$ emission mainly traces the bipolar lobes of M 2-48 and NGC 6537, but their IRAC 8 µm images reveal bright emission at their cores, with faint emission from the bipolar lobes. In sharp contrast, the H$_2$ emission of NGC 650-51 traces its equatorial torus and knotty features in the bipolar lobes, but bright IRAC 8 µm emission also arises from the bipolar lobes.

The lack of Spitzer IRAC observations does not allow us to study the relative spatial distributions of H$_2$ and 8 µm emission for M 1-79 and NGC 6778. As for NGC 6772 and NGC 7048, the radial features of H$_2$ emission seen in M 1-79 are indicative of UV excitation and shielding effects. This may also be the case for the knotty appearance of the equatorial ring of NGC 6778 seen in H$_2$.

### 4.2 Origin of the hydrogen molecular material

The spatial correspondence between H$_2$ and 8 µm emission can be interpreted by the contribution of H$_2$ emission lines into the IRAC 8 µm bandpass. Alternatively, the emission in this IRAC band may be attributed to thermal continuum emission from dust coexisting with molecular hydrogen. The spatial coincidence of molecular hydrogen and dust can have
important consequences for the origin of the molecular material. The dust may act as a shield for H$_2$ molecules from the ionizing UV radiation of the star, and thus dust and H$_2$ may have been present in the nebulae since its formation. Alternatively, the dust grains may act as catalyst for the formation of new H$_2$ molecules on their surface. These phenomena have been studied by Matsuura et al. (2009) for NGC 7293 (Helix Nebula) and van Hoof et al. (2010) for NGC 6720 (Ring Nebula).

The survival of dense, dusty knots, formed during the AGB phase within the ionized zone is critical for the origin of the molecular material (Redman et al. 2003; García-Segura et al. 2006). If the knots are able to survive a long time, then it can be expected that coeval H$_2$ has survived shielded from the UV stellar radiation by the high density and relatively low temperature provided by the knots. However, if these dense knots are rapidly destroyed by UV radiation, then the H$_2$ material detected in old PN should have condensed onto newly formed dust grains.

### 4.3 H$_2$ emission and nebular morphology

The literature provides a wealth of observational evidence supporting the prevalence of H$_2$ emission among bipolar PNe with respect to other morphological types (Kastner et al. 1996, and references therein). Bipolar PNe seem to possess important reservoirs of molecular material, either because they descend from more massive progenitors and have therefore more massive envelopes, or either because the bipolar geometry provides a suitable haven for the survival of molecules in dense equatorial regions. Furthermore, bipolar PNe seem to offer suitable excitation conditions for the excitation and emission of the H$_2$ molecule, may be through shock-excitation, but most likely by offering an appropriate flux of UV exciting photons as the H$_2$ emission is associated in most cases to UV excitation (Likkel et al. 2006).

The quick evolution of the central star of a bipolar planetary nebula—assuming it descends from a massive progenitor—implies it reaches the high effective temperature necessary to provide a suitable UV flux of photons in a short timescale (Aleman & Gruenwald 2004). The combined effects of post-AGB speed evolution and nebular geometry may indeed play an important role, as bipolar PNe that exhibit an equatorial ring structure have much stronger H$_2$ emission than bipolar PNe with a narrow waist (Guerrero et al. 2000).

The prevalence of H$_2$ emission among bipolar PNe led to postulate the so-called Gatley’s rule (Kastner et al. 1996) stating that “the detection of the 2.122 μm S(1) line of H$_2$ is sufficient to determine the bipolar nature of a PN.” This conclusion was based on the correlation between H$_2$ detection and bipolar morphology of a sample of PNe, although some of the PNe in that sample exhibiting H$_2$ emission are not strictly bipolar, i.e., they do not show a butterfly morphology or bipolar lobes connected by an equatorial ring or a waist (Balick 1987; Corradi & Schwarz 1995; Manchado et al. 2000). For example, the physical structure of NGC 6720, the Ring Nebula, has been controverted and sometimes assumed to be bipolar, but a recent study by O’Dell, Sabbadin, & Henney (2007) confirms the closed ellipsoidal shape of the inner shell proposed by Guerrero, Manchado, & Chu (1997).

The detection of H$_2$ emission from PNe in our sample with shell-like morphologies (A 66, M 2-51, NGC 6563, NGC 6772, and NGC 7048) seems to violate Gatley’s rule. We concur that some of these PNe can be described as elipsoidal shells with bipolar extensions (e.g., M 1-79), or barrel-like structures with shorter extensions or annae (e.g., NGC 6563 and NGC 7048). The kinematical information available for some of them in the literature or in the SPM catalogue (López et al. 2012), however, implies that M 2-51, NGC 6563, NGC 6772, and NGC 7048 cannot be described by no means as bipolar PNe. Similarly, the detection of H$_2$
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5 SUMMARY

We have compared the emission detected in IRAC $8 \mu m$ and near-IR $H_2$ images to investigate the nature of the emission observed in this mid-IR IRAC band in a sample of PNe. We confirm that a significant fraction of the IRAC $8 \mu m$ emission can be attributed to $H_2$ line emission, thus revealing the molecular nature of the material seen in these IRAC images. The $H_2$ emission arises from inner shells and outer envelopes or haloes of round and elliptical PNe, as well as from bipolar lobes and dense knots in the equatorial rings of bipolar PNe. We found that $H_2$ emission is not exclusively associated to bipolar PNe, but objects with a barrel-like physical structure and their haloes have also important amounts of molecular hydrogen. We also suggest that the $H_2$ emission from equatorial rings of bipolar PNe arises from discrete knots, rather than from a photo-dissociation region just exterior to the ionized ring.

ACKNOWLEDGMENTS

RAML acknowledges support from CONACYT by the CVU 79367 programs “Becas Nacionales” and “Becas Mixtas de Movilidad en el Extranjero”. He also acknowledges the Instituto de Astrofisica de Andalucia for its great hospitality and the facilities provided for the realization of this work. GRL acknowledges support from CONACYT (grant 177864) and PROMEP (Mexico), RV, MAG, and GRL thank support by grant IN109509 (PAPIIT-DGAPA-UNAM). MAG and GRL acknowledges partial support of the Spanish grants AYA 2008-01934 and AYA 2011-29754-C03-02 of the Spanish Ministerio de Ciencia e Innovación (MICINN) and Ministerio de Economía y Competitividad (MEC) which includes FEDER funds.

Paper based in part on ground-based observations from: the Observatorio Astronómico Nacional at the Sierra de San Pedro Mártir (OAN-SPM), which is a national facility operated by the Instituto de Astronomía of the Universidad Nacional Autónoma de México; the Italian Telescopio Nazionale Galileo (TNG) operated on the island of La Palma by the Fundación Galileo Galilei of the INAF (Istituto Nazionale di Astrofisica) at the Spanish Observatorio del Roque de los Muchachos of the Instituto de Astrofísica de Canarias; the William Herschel Telescope, operated on the island of La Palma by the Isaac Newton Group in the Spanish Observatorio del Roque de los Muchachos of the Instituto de Astrofisica de Canarias; the Nordic Optical Tele-

Figure 11. Spitzer IRS SL spectra of peripheral region of M2-51 (top) and main nebula region (center), and bipolar lobes (low) of NGC 6537. The spectra show the IRAC 5.8 $\mu m$ (green) and 8 $\mu m$ (red) response profiles. Multiple $H_2$ and ionic lines, and some PAHs bands are marked. There is no definite identification for the bright feature at $\sim 7.6 \mu m$, especially in the spectrum of the innermost regions of NGC 6537. We note that this is the wavelength at which the SL1 and SL2 spectra joins, thus we cannot discard it is a spurious artifact.

emission from haloes in A66, NGC 6563, NGC 6772, and NGC 7048 (and probably some more in the literature, e.g., Phillips et al. 2006) does not confirm Gatley’s rule.

Even among the bipolar PNe in our sample, we appreciate notable differences. Sources that do not have an equatorial ring (M2-48 and NGC 6537) show bright 8 $\mu m$ emission at their cores, but the $H_2$ emission arises mostly from the bipolar lobes. On the other hand, the $H_2$ emission from sources with an equatorial ring (NGC 650-51 and NGC 6778) originates from these equatorial regions. In these sources, we note that the $H_2$ emission does not arise from a torus external to the ionized one, but from dense clumps or knots embedded within the ionized ring. This situation is similar to the $H_2$ emission detected in NGC 6720 and NGC 7293 (Hora 2006; Hora et al. 2006; Matsuura et al. 2009; van Hoof et al. 2010; Speck et al. 2002), and reminiscent of the knots that we detect in A66.

Contrary to previous interpretations, the presence of a thick equatorial structure in bipolar PNe does not imply $H_2$ emission: such structures may provide a haven for the survival of hydrogen molecules, but at the same time UV radiation cannot excite these molecules, and thus $H_2$ emission is not produced. Meanwhile the $H_2$ emission from tori of bipolar PNe seems to come from knots that shield themselves from the UV radiation of the central star.
scope, operated on the island of La Palma jointly by Denmark, Finland, Iceland, Norway, and Sweden, in the Spanish Observatorio del Roque de los Muchachos of the Instituto de Astrofísica de Canarias; and the New Technology Telescope at the La Silla Observatory.

This publication makes use of data products from the Two Micron All Sky Survey, which is a joint project of the University of Massachusetts and the Infrared Processing and Analysis Center/California Institute of Technology, funded by the National Aeronautics and Space Administration and the National Science Foundation.

Based in part on photographic data obtained using The UK Schmidt Telescope. The UK Schmidt Telescope was operated by the Royal Observatory Edinburgh, with funding from the UK Science and Engineering Research Council, under a contract with NASA.

This work is based in part on observations made with the Spitzer Space Telescope, which is operated by the Jet Propulsion Laboratory, California Institute of Technology under a contract with NASA.

This publication makes use of data products from the Wide-field Infrared Survey Explorer, which is a joint project of the University of California, Los Angeles, and the Jet Propulsion Laboratory/California Institute of Technology, funded by the National Aeronautics and Space Administration.

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