Variations in the Cyclotron Resonant Scattering Features during 2011 outburst of 4U 0115+63

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ABSTRACT

We study the variations in the Cyclotron Resonant Scattering Feature (CRSF) during 2011 outburst of the high mass X-ray binary 4U 0115+63 using observations performed with Suzaku, RXTE, Swift and INTEGRAL satellites. The wide-band spectral data with low energy coverage allowed us to characterize the broadband continuum and detect the CRSFs. We find that the broadband continuum is adequately described by a combination of a low temperature ($kT \sim 0.8$ keV) blackbody and a power-law with high energy cutoff ($E_{\text{cut}} \sim 5.4$ keV) without the need for a broad Gaussian at $\sim 10$ keV as used in some earlier studies. Though winds from the companion can affect the emission from the neutron star at low energies ($< 3$ keV), the blackbody component shows a significant presence in our continuum model. We report evidence for the possible presence of two independent sets of CRSFs with fundamentals at $\sim 11$ keV and $\sim 15$ keV. These two sets of CRSFs could arise from spatially distinct emitting regions. We also find evidence for variations in the line equivalent widths, with the 11 keV CRSF weakening and the 15 keV line strengthening with decreasing luminosity. Finally, we propose that the reason for the earlier observed anti-correlation of line energy with luminosity could be due to modelling of these two independent line sets ($\sim 11$ keV and $\sim 15$ keV) as a single CRSF.

Key words: X-rays: binaries – pulsars: individual 4U 0115+63

1 INTRODUCTION

4U 0115+63 is a high mass X-ray binary system, first discovered in the Uhuru satellite’s sky survey (Giacconi et al. 1972; Forman et al. 1978), with more than 15 subsequent outbursts recorded till date (Boldin et al. 2013). The system consists of a pulsating neutron star with spin period $\sim 3.61$ s (Cominsky et al. 1978) and a B0.2Ve main sequence star (Johns et al. 1978), with an orbital period of $\sim 24.3$ days (Rappaport et al. 1978). The distance to this binary system has been estimated to be $\sim 7$ kpc (Negueruela & Okazaki 2001). The source exhibits luminous Type II X-ray outburst during which multiple cyclotron resonance scattering features (CRSF) have been observed in the X-ray spectrum, with 5 detected harmonics (Santangelo et al. 1999, Ferrigno et al. 2009). CRSF are caused by scattering of X-ray photons from electrons in the accreting plasma channeled by the NS magnetic field. The energy at which these lines occur is given as $E_{\text{cyc}} = 11.6B_{12} \times (1 + z)^{-1}$ keV (Coburn et al. 2002). Here $B_{12}$ is the local magnetic field (in units of 10$^{12}$ Gauss) and $z$ is the gravitational redshift in line energy. Thus cyclotron lines give us a direct probe of the local magnetic field near the scattering regions.

Cyclotron line parameters of many sources are found to vary with the phase of rotation (see Heindl et al. 2004, for a review), and the varying luminosity of the outburst (Becker et al. 2012). However, the variation in energy of the fundamental CRSF of 4U 0115+63 with luminosity has been the source of some debate. Nakajima et al. (2006), Tsygankov et al. (2007), Li et al. (2012) find an anti-correlation between the line energy and luminosity, whereas Müller et al. (2013) find this anti-correlation to be an artifact of the continuum spectral modelling. Boldin et al.
of these satellites is explained in the following sub-sections. The reduction of data-sets from each observations performed at different luminosity levels are also made when the source was near its peak luminosity. Figure 1 shows the variations in the count rates as measured with MAXI during the 2011 outburst. The symbols represent the pointed observations and the horizontal line shows the luminosity level (4 x 10^{37} ergs s^{-1}) around which the previous observations have found a sharp change in fundamental CRSF energy.

have pointed out the cause of this reported dichotomy to be due to the use of a broad Gaussian like emission feature to model the continuum in some of the works.

In this paper, we study multiple observations during the 2011 outburst of 4U 0115+63 performed with Suzaku, RXTE, Swift and INTEGRAL satellites, providing us with a wide-band coverage from 0.5 keV to ~ 60 keV with high signal-to-noise. As shall be seen in subsequent sections, this availability of wide bandwidth data is very important to correctly model both the continuum and the CRSF in the source spectrum. From the results of our spectral analysis, we find evidence for two sets of cyclotron lines whose parameters vary with source luminosity. In the following sections, we describe the observations and reduction of data that we used in spectral analysis of these data-sets in and the inferences and possible implications in.

2 OBSERVATIONS AND DATA REDUCTION

In all, we analysed data obtained by different X-ray observatories over the 15 days of the 2011 outburst. Table I lists all the RXTE and Suzaku observations which were available in the archive and a Swift and an INTEGRAL observation, made when the source was near its peak luminosity. Figure 1 shows the variations in the count rates as measured with MAXI during the 2011 outburst. The pointed observations performed at different luminosity levels are also marked in this figure. The reduction of data-sets from each of these satellites is explained in the following sub-sections.

Table 1. A summary of pointed observations during the 2011 outburst of 4U 0115+63.

| MJD       | Instrument | ObsId   | MAXI$^a$ | Exposure$^b$ |
|-----------|------------|---------|----------|--------------|
| 55730.06  | RXTE       | 96032-01-01-00 | 0.788  | 6.53         |
| 55736.34  | Swift      | 00031172010  | 0.925  | 6.69         |
| 55736.34  | INTEGRAL   | 106100650010 | 0.925  | 2.13         |
| 55738.87  | RXTE       | 96032-01-02-00 | 0.993  | 4.52         |
| 55739.51  | RXTE       | 96032-01-02-01 | 0.745  | 3.69         |
| 55743.91  | RXTE       | 96032-01-03-00 | 0.682  | 5.02         |
| 55746.90  | RXTE       | 96032-01-03-02 | 0.526  | 9.04         |
| 55747.53  | Suzaku     | 4060475000  | 0.526  | 24.27        |
| 55750.82  | RXTE       | 96032-01-04-00 | 0.352  | 16.91        |
| 55750.82  | Suzaku     | 4060475000  | 0.352  | 81.08        |
| 55751.32  | RXTE       | 96032-01-04-02 | 0.388  | 2.08         |
| 55753.00  | RXTE       | 96032-01-04-03 | 0.248  | 1.85         |
| 55753.42  | RXTE       | 96032-01-04-04 | 0.238  | 0.73         |

$^a$MAXI rate (in counts cm^{-2} s^{-1}) is of nearest observation.
$^b$Exposure time is in kilo-seconds.

2.1 RXTE Observations

Among the RXTE observations, we used all but two of them. Obs-Id 96032-01-04-03 and 96032-01-04-04 were made when source flux was very low, and with short exposure times (1856s and 736s respectively). Due to poor count statistics, we were unable to constrain the CRSF parameters for these observations, making them unsuitable for the present work. RXTE spectra were obtained from the raw data files using FTOOLS from HEASoft v 6.15.1. We used data from both the HEXTE and the PCA detectors. PCU2 Science Array data were used for generating PCA spectrum, and Cluster A Science Array data were used for HEXTE spectrum. HEXTE background was obtained from Cluster B, and dead-time corrections were applied to both the source and background HEXTE spectra. For PCA, we found the dead-time to be a maximum of 5%, which lowered its flux by about 5%. We did separately correct for this. Data grouping and usage of systematic errors for the RXTE spectra are shown in Table 2. In all, the RXTE provided usable data covering the 3 keV to 50 keV band.

2.2 Suzaku Observations

Suzaku had two long duration observations made during the course of the 2011 outburst. For reduction of Suzaku data, we used FTOOLS from HEASoft 6.15.1 with CALDB updated till June 2014. We used data from the XIS (0.6 keV - 10 keV) and PIN (15 keV - 60 keV). The GSO data had very low SNR and hence was not used. Although the PIN does collect data from 12 keV onwards, we discarded the data from 12 keV to 15 keV due to high uncertainty in PIN background. Suzaku data above 10 keV and PIN data above 60 keV were too noisy to be of use for our analysis. As a result we were not able to effectively cover the energy range (10 keV - 15 keV), which is important to model the fundamental cyclotron line in this system. To overcome this, we used simultaneously taken

1 see Suzaku Data Reduction ABC guide (version 4.0)
RXTE observations in conjunction with our Suzaku data-sets. Since Obs-Id 406049010 (Suzaku) and Obs-Id 96032-01-04-00 (RXTE) had some overlap, we used the portion where overlap existed. This was done by creating a GTI which covered the common interval between the RXTE and Suzaku GTI files. However Obs-Id 406048010 (Suzaku) had no overlap with any other RXTE observation. Thus, we did not use this data-set.

The XIS data were taken in the 1/4 windowed mode (of size 256 pixels) at normal clocking speeds to reduce the effect of pile-up. We found significant pile-up up to a maximum level of ~17% in the central regions and corrected for it using the recipe of John Davis.

This was achieved by rejecting a central circular region of size ~25' from the XIS image for which the computed Pile-up percentage was greater than 6%. After extraction of data from individual XIS chips, the XIS 0 and 3 data were combined using the addascaspec tool. This resulted in two sets of spectral files; one from the back illuminated (BI) CCD and the other from the combined front illuminated (FI) CCDs. XIS data below 0.8 keV were rejected owing to discrepancies between the FI and BI spectra at a ~3σ level. A mismatch at a ~2σ level between the two spectra around the Si K-edge was also noticed, but we did not discard this part of the spectrum as it was important for fitting our continuum model. This led to poorer values of the χ² statistic that we use for fitting the spectral data.

The PIN data were extracted using the tool hxdpinxbp1. In the joint analysis of Suzaku and RXTE, we required a cross-normalisation of ~1.57, between the Suzaku XIS and PIN, much larger than the recommended value of 1.16 given in the Suzaku Data Reduction ABC guide (version 4.0). Therefore, we investigated this issue in some detail. To verify the PIN data, we extracted individual 64 PIN count spectrum as explained in the Suzaku note on estimating PIN noise and found no discrepancies. We also extracted night-earth data for the PIN observations to compare and see if the estimated backgrounds were correct. Fig. 2 compares the night-earth and the “tuned” background spectral data. The two background spectra are similar in shape and in flux.

We further examined the validity of the background subtraction by comparing the net (source minus background) and background light curves. The lack of correlation amongst the two indicates that the PIN background was estimated correctly. Finally we checked if the XIS spectra were properly area corrected. The XIS data were taken in the 1/4 window mode with normal clocking, thereby making a rectangular source footprint. By choosing both rectangular and circular source regions for extraction we found no difference in the angular source footprint. By choosing both rectangular and circular source regions for extraction we found no difference in the angular source footprint. By choosing both rectangular and circular source regions for extraction we found no difference in the angular source footprint.

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normalisation factor of 1.57 for all our subsequent models and fits (also see Sect. 3).

2 http://space.mit.edu/CXC/software/suzaku/
3 http://heasarc.gsfc.nasa.gov/docs/suzaku/analysis/pinnoise.html
4 http://www.swift.ac.uk/analysis/xrt/rmfs.php

2.3 Swift and INTEGRAL observation

We selected this particular observation from multiple Swift observations of the source, because it was simultaneously taken with an INTEGRAL observation (thereby enabling wide-band spectral coverage) and it was near the point having highest MAXI counts. This observation gave us a wide-band data-set near the peak luminosity with similar energy coverage to the Suzaku data-set at much lower luminosity levels. However, as pointed out in section 3, the low effective area of INTEGRAL JEM-X resulted in lower signal to noise in the energy band from 11 keV to 17 keV , than for the Suzaku data-set.

Swift data reduction was carried out using FTOOLS from HEASoft 6.15.1 with CALDB updated till June 2014. INTEGRAL analysis was done using DSA 10.0 with its calibration files also updated till June 2014. We used the Swift XRT data from 0.5 keV to 9.7 keV for our spectral analysis. These data were taken in the windowed timing (WT) mode of XRT. The maximum observed count rate was about 40 counts s⁻¹. This ensured that the XRT data was not piled up (see Romano et al. 2006). The XRT data from 0.4 keV to 1 keV did show an excess due to uncertainty in the response modelling of XRT WT mode, but upon using the position dependent WT response files, this apparent excess was removed. Hence we used these position dependant response files for all our analysis.

We used two instruments - the JEM-X and the IBIS from the INTEGRAL’s suite. The JEM-X 2 data (from JEM-X) and the ISGRI data (from IBIS) were extracted with higher spectral binning than used in the standard pipeline. The JEM-X data were extracted from 3 keV to 30 keV and the IBIS data from 15 keV to 100 keV , as per the instructions in the IBIS Analysis User Manual (Issue 10.0) . The latest calibration files were used for generating the spectra and the rebinned rmfs. After extraction, the data were grouped and systematic errors added as given in Table 2. We used spin phase averaged spectra in all the observations. For modelling the spectra, we used XSPEC v12.8.1g (Arnaud 1996).

3 SPECTRAL ANALYSIS

We began our spectral analysis with the wide-band Suzaku dataset (obs-id 406049010) taken alongwith RXTE dataset (obs-id 96032-01-04-00). For other datasets, we used the results from this analysis as a template because this data-set had both wide-band spectral coverage and high signal to noise. As noted previously, when we held the PIN-XIS cross-normalisation fixed at 1.16 we found a significant offset of only the PIN spectrum from the other instrument spectrum in the ratio plots (refer Fig. 3 for the best fit cutoffpl based model with CRSFs). When we let the normalisation parameter free, we found that it gave a best fit value of 1.57 which we adopted for all our subsequent analysis.
To model this wide-band spectrum, we started off by using the simple cutoffpl based continuum modified by interstellar absorption and CRSFs. At low energies ( < 2 keV), effects of local absorption and emission features from the plasma wind of the companion star are known to affect the spectrum (Suchy 2011). However, we did not specifically account for this in the spectrum. Complete modelling of the wind spectrum is preferably done with higher spectral resolution data from gratings. As mentioned in Sec. 2.2, the Suzaku data-set had calibration uncertainties around the detector’s Si K-edge (from \( \sim 1.6 \) keV to \( \sim 2.5 \) keV). Thus modelling the wind effects was difficult with this data-set. As these effects are restricted to lower energies, and as we do not consider data below 0.8 keV in our analysis, we do not expect this to change our continuum model estimates significantly.

The interstellar absorption was modelled by an updated version of tbabs\(^5\) using abundances of Wilms et al. (2000) and cross-sections as given in Verner & Yakovlev (1995).

This model fit well if we took data above 3 keV, but showed a significant excess when we included the data from

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\(^{5}\) http://pulsar.sternwarte.uni-erlangen.de/wilms/research/tbabs
0.8 keV to 3 keV. This is clearly seen in the top left panel of Figure 4, where we obtained \( \chi^2/\text{dof} = 2018.60/474 \). On using a blackbody (\texttt{bbody}) to account for this excess, we found that the spectral data were well fitted to give a significant improvement in fit \( \chi^2/\text{dof} = 731/472 \). The temperature and radius of the blackbody so obtained are listed in Table 3. The radius was computed from the normalisation of the blackbody model \texttt{bbody}, which depends on the luminosity and distance to source. The distance was taken as 7 kpc and luminosity was taken as \( L_{\text{bb}} = \sigma \cdot T^4 \).

Table 3. Blackbody radius and area for different observations.
Note that \textit{RXTE} observations have higher errors.

| Instrument | Day (MJD) | \( \text{KT (keV)} \) | Radius (kms) |
|------------|-----------|----------------------|--------------|
| \textit{RXTE} | 55730.06 | 1.08 ±0.18 | 13.12 ±3.78 |
| \textit{Swift / INTEGRAL} | 55736.34 | 0.93 ±0.19 | 17.39 ±4.31 |
| \textit{RXTE} | 55793.51 | 0.40 ±0.41 | 14.69 ±0.9 |
| \textit{RXTE} | 55743.91 | 0.79 ±0.20 | 19.57 ±0.47 |
| \textit{RXTE} | 55747.89 | 0.78 ±0.00 | 18.83 ±0.00 |
| \textit{Suzaku / RXTE} | 55750.82 | 0.73 ±0.01 | 13.95 ±0.41 |
| \textit{RXTE} | 55751.32 | 0.92 ±0.26 | 10.54 ±0.37 |

\( ^a \) All errors in this table are 1\( \sigma \) deviations
\( ^b \) calculated as \( \sqrt{\text{area}} \)
\( ^c \) Could not put error constraints on blackbody temperature

We tried using \texttt{comptt} model to account for this low energy excess, as indeed was tried in \cite{Ferrigno2009}. This gave an improvement in fit with the \( \chi^2 \) reducing from 688/469 to 583/466. This large improvement strongly indicates that the blackbody seed photons are indeed comptonized. The spectral fits for \texttt{comptt} were obtained though, for a electron plasma temperature (\( kT_e \)) of 1.2 keV, a seed photon temperature of 0.2 keV and plasma optical depth (\( \tau \)) of 35. The \texttt{comptt} code itself is meant for use for plasma temperature greater than 2 keV and at such high \( \tau \), the shape of the spectrum is very similar to the Wien tail of a blackbody. Using the \texttt{comptt} model for this Wien tail alone, also lead to an increase in the value of the absorption column from \( N_H = 1.3 \times 10^{22} \text{ cm}^{-2} \) to \( N_H = 1.7 \times 10^{22} \text{ cm}^{-2} \). When we tried the same model in the \textit{Swift / INTEGRAL} spectrum, we found a very small reduction in the \( \chi^2 \) from 146/166 to 145/163. Furthermore, none of the \textit{RXTE} observations gave an improvement in fit \( \chi^2 \) when using the \texttt{comptt} model as compared to the \texttt{bbody} model. Finally, we found very little differences in the CRSF energies, while using the simple blackbody model as compared to the comptonized blackbody. Since the data that we analysed could not help us pin down the nature of the low energy continuum, we used the simple blackbody for all subsequent analysis. However, we do note that probing this component further may be important to understand the overall continuum, which in turn may affect the cyclotron line results.

As stated in \cite{Suchy2011} and \cite{Muller2013}, we found the column density (\( N_H \)) to be strongly correlated with the power-law index (\( \Gamma \)) and to be varying across observations. To prevent any unwanted effects due to changing \( N_H \) on our fits, we fixed the column density for all observations to the best fit value of \( N_H = 1.3 \times 10^{22} \text{ cm}^{-2} \) obtained from the wide-band \textit{Suzaku} observation. We also used a narrow emission line with its centroid fixed at 6.4 keV and width fixed at \( 10^{-4} \) keV to model the Fe K-\( \alpha \) fluorescence emission which has been observed in many other HMXB NS binary systems and is also expected to be seen in 4U 0115+63 (\cite{Torrente2011, Muller2013}). For the wide-band \textit{Suzaku} observation, this narrow Fe-line component gave an improvement in \( \chi^2/\text{dof} \) from 756/473 to 731/472. Upon testing this with the \texttt{simtest} script of \texttt{XSPEC}, we found it to be significant to a level greater than 3\( \sigma \). This test was done using XIS data only, to isolate any effects due to viewing geometry and asymmetric emission

We tried next to model the high luminosity wide-band observations. Using the \textit{Swift / INTEGRAL} spectrum, we found cyclotron line variation in 4U0115+63 (\cite{Mihara1990}), with each

\[ N_H (10^{22} \text{ cm}^{-2}) \]
\[ \Gamma = 2.16 \pm0.48 \]
\[ E_{\text{cut}} \text{ keV} = 4.18 \pm0.12 \]
\[ kT_{\text{bb}} \text{ keV} = 0.94 \pm0.31 \]
\[ b_{\text{norm}} (x 10^{-3}) = 15.5 \pm0.7 \]
\[ \text{Fe line eq.width (eV)} = 11.7 \pm0.7 +20.5 \pm+3.4 \]
were able to fit the spectrum from 0.5 keV up to 60 keV using only the cutofpl continuum with absorption fixed at $N_{\text{H}} = 1.3 \times 10^{22} \text{ cm}^{-2}$ and CRSFs ($\sim 11$ keV, 20 keV and 33 keV) to get a $\chi^2/dof = 172.8/168$. On adding the blackbody component, this improved to $\chi^2/dof = 146.1/166$. We tested the significance of this using the lrt script of XSPEC. This gave a significance of greater than 3σ for the blackbody model for 408 iteration runs. When taken alongside the case of Suzaku / RXTE data-set, we see that the presence of a soft X-ray component is justified. As stated above, modeling this component is important. For the sake of consistency, we use the blackbody model, as using the comptt model for the soft X-ray component did not give us any improvement over the blackbody based model for this observation. While using this blackbody and cutofpl continuum, we could not fit this data-set with the 5th line as earlier and the fit was consistent with only 3 CRSFs. A summary of the continuum parameters obtained for the two wide-band data-sets is given in Table 5. Errors, unless otherwise mentioned are quoted for a level of 90% confidence in all cases.

We finally analysed the set of standalone RXTE observations. Due to the lack of coverage of the low energy bands in these data-sets, it was difficult to constrain the blackbody parameters. Hence we fit the blackbody by starting from a guess value based on the fit parameters obtained from the wide-band data-set and let the fit routine converge to give the values as quoted in Table 3. The blackbody temperature was frozen to this best fit value while computing errors on other fit parameters. Excepting one, all the RXTE data-sets gave an improvement in the fit statistic when using 4 CRSFs as compared to the fit with 3 CRSFs (see Table 4). The exception was the first data-set, taken during the rising phase, which needed only two CRSFs to describe the spectrum. We consider this case separately in the Sec. 4. Given this improvement when using 4 CRSFs, we decided to test its statistical significance for describing the spectral data.

3.1 Statistical significance of the 4th cyclotron line
Computing detection significance of CRSFs must be handled differently from the standard techniques usually employed for emission features. First, following the method outlined in Orlandini et al. (2012) we construct the F-statistic directly as a ratio of the normalised fit $\chi^2$ (Orlandini et al. 2012; Bevington & Robinson 1992), as against using the standard F-test implemented in XSPEC, which constructs the F-stat as a ratio of change in normalised $\chi^2$ to the original $\chi^2$.
Table 4. Best-fit parameters for the CRSFs.

| Day (MJD) | \( E_{\text{cyc}} \) (keV) | \( \text{FWHM} \) (keV) | \( \tau \) | \( \chi^2/\text{dof} \) | Eq. width (keV) | \( \Delta\chi^2 / \Delta \text{dof} \) | F-stat % |
|-----------|-----------------|-----------------|---|----------------|----------------|----------------|--------|
| 55730.06  | 16.33 \(+0.32\) | 9.38 \(+0.76\) | 1.43 \(+0.13\) | 57 / 73 | 11.98 \(+0.72\) | 0 / 0 | 0 |
| (RXTE)    | 37.29 \(+1.15\) | 4.03 \(+0.56\) | 0.45 \(+0.14\) | 2.38b |
| 55736.34  | 10.31 \(+1.58\) | 5.96 \(+4.58\) | 0.97 \(+0.31\) | 146 / 166 | 6.21 \(+4.23\) | 0 / 0 | 0 |
| (Swift / INTEGRAL) | 20.69 \(+5.21\) | 9.92 \(+5.24\) | 1.26 \(+0.31\) | 11.98 \(+4.50\) |
| 35.75 \(+2.19\) | 4.00 \(+0.9\) | 0.45 \(+0.36\) | 2.38b |
| 55738.87  | 11.71 \(+0.37\) | 4.33 \(+1.22\) | 0.62 \(+0.25\) | 47 / 67 | 3.35 \(+1.18\) | 14.91 / 3 | 83 |
| (RXTE)    | 15.82 \(+0.68\) | 1.1 \(+1.66\) | 0.13 \(+0.31\) | 10.48 \(+1.59\) |
| 19.86 \(+0.01\) | 8.82 \(+1.91\) | 1.21 \(+0.38\) | 0.17 \(+0.19\) |
| 35.52 \(+0.65\) | 4.00 \(+0.9\) | 0.35 \(+0.16\) | 1.9b |
| 55739.51  | 11.66 \(+0.89\) | 5.12 \(+3.7\) | 0.74 \(+0.2\) | 43 / 67 | 4.49 \(+4.56\) | 5.81 / 3 | 64 |
| (RXTE)    | 16.08 \(+2.16\) | 2.09 \(+2.05\) | 0.14 \(+0.1\) | 0.42 \(+1.40\) |
| 20.42 \(+2.17\) | 8.52 \(+2.25\) | 1.31 \(+0.31\) | 10.58 \(+3.59\) |
| 35.24 \(+1.38\) | 4.00 \(+0.9\) | 0.38 \(+0.14\) | 2.6b |
| 55743.91  | 11.25 \(+0.98\) | 4.13 \(+3.04\) | 0.45 \(+0.33\) | 31 / 67 | 2.50 \(+2.94\) | 10.17 / 3 | 84 |
| (RXTE)    | 15.2 \(+1.93\) | 2.82 \(+2.76\) | 0.38 \(+0.27\) | 1.44 \(+2.35\) |
| 19.56 \(+1.97\) | 7.57 \(+1.73\) | 1.17 \(+0.42\) | 8.86 \(+2.09\) |
| 32.77 \(+0.85\) | 4.00 \(+0.94\) | 0.63 \(+0.15\) | 2.9b |
| 55747.59  | 11.26 \(+0.86\) | 3.53 \(+4.04\) | 0.33 \(+0.39\) | 52 / 67 | 1.63 \(+1.11\) | 14.99 / 3 | 81 |
| (RXTE)    | 14.93 \(+0.83\) | 2.76 \(+2.41\) | 0.35 \(+0.51\) | 1.33 \(+1.40\) |
| 18.01 \(+0.98\) | 9.42 \(+1.73\) | 1.33 \(+0.32\) | 11.67 \(+1.34\) |
| 33.79 \(+1.56\) | 4.00 \(+0.9\) | 0.26 \(+0.14\) | 1.4b |
| 55750.82  | 10.77 \(+0.32\) | 0.96 \(+0.56\) | 0.12 \(+0.03\) | 688 / 469 | 0.17 \(+0.10\) | 42.68 / 3 | 74 |
| (RXTE / Suzaku) | 14.39 \(+0.33\) | 4.19 \(+0.92\) | 0.69 \(+0.20\) | 3.48 \(+1.60\) |
| 19.53 \(+0.33\) | 8.33 \(+2.52\) | 1.13 \(+0.19\) | 9.52 \(+2.21\) |
| 31.35 \(+0.98\) | 4.00 \(+0.9\) | 0.56 \(+0.16\) | 2.8b |
| 55751.32  | 10.70 \(+0.0\) | 2.1 \(+2.1\) | 0.12 \(+0.09\) | 57 / 67 | 0.37 \(+0.59\) | 4.32 / 3 | 55 |
| (RXTE)    | 14.51 \(+0.95\) | 4.67 \(+3.08\) | 0.73 \(+0.28\) | 4.06 \(+1.07\) |
| 19.55 \(+0.0\) | 8.7 \(+1.52\) | 1.34 \(+0.25\) | 10.90 \(+1.84\) |
| 31.98 \(+1.16\) | 4.00 \(+0.9\) | 1.05 \(+0.42\) | 4.4b |

\( ^a \) Change / reduction in \( \chi^2 \) on addition of extra CRSF / change in number of dof - denotes the confidence that a 4 CRSF model is better than a 3 CRSF model

\( ^b \) Cannot quote error on this, as width is frozen

We try to compute the significance of the 4th CRSF at 11 keV, which we used as stated previously to improve the fit. For the RXTE/Suzaku joint spectral fit, the above method gave us an F-statistic value of 1.06 with a probability value for occurrence due to random noise (PCI) equal to 0.26. We also noted that ignoring data from 1.6 to 2.5 keV (the energy range with calibration mis-match between the two XIS detectors) reduces this probability to 0.24 with very little change in spectral parameters. This is not enough to claim detection of a fourth cyclotron line. The F-stat based test performed above, looks at what fraction of the variance of the data-set can be explained by the model. This test looks at the percentage of data variance that a new model (with all its components) can explain as compared to the percentage of data variance explained by the old model (with all its components). Unaccounted residuals exist in the soft part of our fit spectrum below 3 keV. This might account for the low reduction in percentage variance accounted for by the addition of a cyclotron line at the higher energies (11 keV). We checked each of the other RXTE observations individually with this F-test (results listed in Table 1). The broad-band continuum is not completely sampled by the RXTE data alone and this could explain the low improvement in \( \chi^2 \) for the individual RXTE data-sets. To improve the statistics, we did a joint fitting of all the data-sets using both the 3 CRSF and the 4 CRSF models. In doing this fit, we let the line energies for each observation be independently estimated by leaving them untied. This gave an F-statistic value of 1.07 and a probability of chance occurrence of 0.15 (or a confidence level of 85%).

Finally, we tried a more robust numerical evaluation of significance from Monte-Carlo simulations. We did this by finding the probability of false detection of the 4th CRSF, assuming the 3 CRSF model to be true. By simulating spectra using the XSPEC \texttt{lrt} script, we generated a large number (7438) of simulated data-sets following the continuum model with 3 CRSFs modified by statistical noise. The \texttt{lrt} script generates these data-sets by using the fit covariance matrix to make a random draw of the fit parameters. The model so obtained, is convolved with the response matrices of the individual detectors, and statistical (Poisson) noise is added to each such simulated data-set. We searched for the pres-
ence of the 4th cyclotron line in each of these data-sets by trying to fit them with the continuum and 4 cyclotron lines, and compared it to a fit with the continuum and 3 cyclotron lines. We tabulated the \( \chi^2 \) fit value for each such effort. To check if the observed 4th CRSF was significant, we compare the statistic of each of the simulation runs against the statistic of the observed data. The form of the statistic we used was the F-stat \( (\text{fst}) \) constructed as mentioned before (see also Sartore et al. [2013]). We obtained three instances where the simulated \( \text{fst} \) was as high as the observed \( \text{fst} \). However, in none of these instances, the 4th CRSF fit with centroid near 11 keV. This gave a PCI of 3 in 7438, or a significance of 3.5\( \sigma \) for the observed CRSF at 11 keV.

This result, though needs to be treated with care because of the poor \( \chi^2 \) of the original fit \( (731/472) \) vs \( (688/469) \). We proceed assuming the 4 CRSF model to be the better one, but only after noting that the poor initial fit could influence the Monte-Carlo results.

We then looked at the validity of the 4 CRSF model in describing other data-sets. All the individual RXTE spectra do not have wide-band coverage and have lower spectral resolution than say the combined Suzaku / RXTE data-set. This could be a reason for the marginal improvement in the \( \chi^2 \) statistic. The fact that multiple observations showed signs of 4 CRSF features, with an improvement in \( \chi^2 \) would lead to an increased relevance of this detection. Additionally, given that each of these observations gave nearly similar centroid energies for all 4 CRSFs (see Fig. 5) would also increase the significance of our detection.

The possible reasons for two observations not requiring 4 CRSFs to describe their spectra are discussed herewith. The Swift and INTEGRAL simultaneous observations not fitting with the 4 CRSF model could be either because the line at 15 keV does not exist in this observation, or because the line is too weak to be detected by the JEM-X detector due to its much lower collection area and poorer signal to noise than the RXTE/PCA. This is where we find the high signal to noise of the combined Suzaku / RXTE data to be critically important. The RXTE standalone observation which did not show 4 CRSFs, though had a high signal to noise all through the expected CRSF energy ranges. We take the detection of only 2 CRSFs in this observation to be a valid result and discuss the possible reasons for non detection in this data-set in the next section.

A useful indicator for evaluating the strength of the absorption lines, and to see if they are physically relevant is the line equivalent width. We computed the equivalent width as

\[
EW \text{ (keV)} = \int_{E_1}^{E_2} (1 - e^{-\tau P(E)}) \, dE
\]

\[P(E) = \frac{(W \frac{E}{E_{\text{cyg}}})^2}{(E - E_{\text{cyg}})^2 + W^2}
\]

where \( P(E) \) is the line profile that we use, \( W \) is FWHM of the line, \( E_{\text{cyg}} \) is its energy and \( \tau \) is its depth. By simple error propagation of the variances obtained from the fit covariance matrix, we can get an idea of errors on the estimate. We calculated this for all 4 CRSFs to check for their relevance. The results of such calculations are plotted in Fig. 5. We try to make sense of these results in the next section.

\[4\] DISCUSSION

4.1 Summary of results

In this work we have analysed the spectra from different X-ray observatories covering the 2011 outburst of the HMXB 4U 0115+63. These include eight observations spread across twenty days including two wide-band observations compris-
The two major results from our analysis were inconclusive. However, such Gaussian features have previously been attributed to possible cyclotron emission not required for our analysis. Such Gaussian features have been found to be statistically significant in both of the broadband observations, as well as the standalone RXTE observations. Excess emission features in the shape of a Gaussian, as reported in previous works e.g. Müller et al. (2013) was not required for our analysis. Such Gaussian features have previously been attributed to possible cyclotron emission (Becker & Wolff 2007; Ferrigno et al. 2014), however such results were inconclusive.

(ii) Two sets of cyclotron lines: From the spectral analysis we find evidence for presence of two independent sets of cyclotron lines, each with a harmonic, with fundamental energies centred at \( \sim 11 \) keV and \( \sim 15 \) keV line. Previous works find either one of the two fundamental lines, with the 15 keV line found at lower luminosities, and the 11 keV line found at higher luminosities (Tsygankov et al. 2007), resulting in the inference of anti-correlation of the line energy and luminosity. To check whether the 4 cyclotron lines are two independent sets of harmonics we compared the energy for the fundamental 11 keV line.

Figure 7. Trends in the fundamental lines. Top panel shows the effect of not using 4 CRSF lines, which leads to the anti-correlation with luminosity as reported in previous works. Bottom panel shows the change in line equivalent width of the fundamental of each set of the 4 CRSF model. The vertical line is placed at the same luminosity level as in Fig 1. The rising phase data-set, which does not seem to follow the trend is marked separately. See text for details.

(a) There is no luminosity dependence of the fundamental cyclotron line at 11 keV if we use the 4 CRSF model, where both the \( \sim 11 \) and \( \sim 15 \) keV lines are present. The anti-correlation appears if the 3 CRSF model is used, as shown in the top panel of Fig. 5. This could be an artifact resulting from incorrect spectral modelling with only 3 cyclotron lines, instead of 4 CRSFs, as we discuss below.

(b) In the 4 CRSF model, the line equivalent widths of the two lines at 15 keV and 11 keV show opposite variations with luminosity as seen in the bottom panel of Figure 7.

When only 3 cyclotron lines are used to model the spectra, the two lines at \( \sim 11 \) keV and \( \sim 15 \) keV are modelled by a single CRSF component. Since the equivalent widths of these two lines change with luminosity, the single averaged CRSF component is closer to the line with higher equivalent width in the given observation. Simulations have shown that often the second harmonic is deeper and more prominent (Araya-Góchez & Harding 2000; Schönherr et al. 2006) due to photon filling and emission by de-excitation near the fundamental CRSF energy. Results from previous outbursts (Li et al. 2012; Boldin et al. 2013) and our analysis demonstrate the near constant line energy of the 20 keV line. This result too would lead us to expect a near constant line energy for the fundamental 11 keV line.

If we split the 4 CRSFs that we obtained into two sets of harmonics with one at (11 keV and 20 keV ) and the other at (15 keV and 33 keV ), it helps us explain our observations as listed below:

(a) As shown, this explains the reason for observations of an anti-correlated fundamental CRSF with luminosity.

(b) It could be the reason why this source is the only one to have shown multiple (upto 5) harmonics of the fundamental cyclotron line in its spectrum. If we have two such line forming regions with fundamentals at 11 keV and 15 keV , and each region showed the presence of 2 harmonics, then they can easily be confused for multiple harmonics from a single 11 keV fundamental CRSF.

(c) This would explain why the CRSFs we obtain at 20 keV varies so little, whereas the one we get at 33 keV varies a lot more. Under our hypothesis, the \( \sim 33 \) keV line in our observations would be a combination of the 2nd harmonic from the 15 keV set (at \( \sim 30 \) keV ) and the 3rd harmonic from the 11 keV set (at \( \sim 33 \) keV ). As seen from Fig. 6 and 7 the \( \sim 33 \) keV line shifts to higher energies when the \( \sim 11 \) keV line becomes stronger.

(d) Finally, it gives a plausible reason for detection of only two CRSFs in the RXTE observation taken during the rising phase. If, for some reason, the 11 keV line set is either not present or has very weak signatures in the rising phase of the outburst, then the observed CRSF energies in the rising phase would correspond to the 15 keV line set only. The values we get for the observed CRSFs seem very close to this.

4.2 Possible origin of the cyclotron line sets

In this section we discuss where the two scattering regions could be located and what causes the CRSF strengths from these regions to vary in such a manner.
The first possibility is emission from two different regions at different heights on the same pole. If we assume a dipole like magnetic field structure of the Neutron Star (NS) the difference of $\sim 4$ keV between two sets of lines can be caused if the difference in height of the emitting regions is $\sim 1.1$ km, which is approximately the shock height above the neutron star surface (Becker et al. 2012). If the 15 keV line set originates in the fan beam at the base of the mound, and the 11 keV line set in the pencil beam from the top of the shock, then variation in line equivalent widths with luminosity can be explained in terms of varying strengths of the fan and pencil beam emissions. At higher luminosities, the fan beam is expected to dominate, whereas both beams will be visible at intermediate luminosities levels (see Fig. 1 of Becker et al. 2012). However, the effect of such variations of luminosity on the equivalent width is unclear and not adequately addressed by existing theoretical models.

The second possibility is emission from different poles. Sasaki et al. (2012) have analysed the pulse profiles of this source and found a $\sim 60^\circ$ offset between the position of one of the poles and the antipodal position of the other pole. This non dipolar field structure could lead to different local fields at the scattering regions of each pole, which in turn could cause the difference in energies of the lines formed in each of them. The variation in accretion rate onto one pole as compared to the other can cause the variation in the equivalent widths as seen in bottom panel of Fig. 7.

We can get further indicators by looking at the pulse resolved energy spectrum and pulse phase-lag spectrum. We examined the energy dependant phase lag of the pulses as computed by Ferrigno et al. (2011) for the observation with simultaneous RXTE/Suzaku data. We have computed the phase lags using the data from the higher time resolution RXTE PCA detector, in a manner similar to that of Ferrigno et al. (2011). As seen in Figure 8, the pulse has two distinct peaks. We obtain the phase lags by taking phases corresponding to the main peak at a reference energy and cross-correlating them against the same set of phase bins at other energies. We have performed this for both the main and secondary peaks, with phase bins taken as shown in bottom panel of Fig. 8. In Ferrigno et al. (2011), the authors report that the most negative phase shifts occur at energies near the CRSF energies. This has been attributed to a change in beam pattern at the CRSF energies, with the photons at the CRSF band lying along a pencil beam whereas the rest of the emission is dominated by a fan beam.

In our analysis, we found that main peak’s phase-lag spectrum had a similar set of minima at $\sim 11$ keV, $\sim 23$ keV and $\sim 39$ keV. However, the secondary peak’s phase-lag spectrum had minima at $\sim 16$ keV and $\sim 30$ keV. The low count rate and increasing errors for energies above 45 keV made it difficult for us to find these dips at higher energies. The phase bins for performing the cross-correlation were chosen in such a manner that the pulse profile at the reference energy band (taken as the band between 8.17 keV to 10.63 keV, similar to Ferrigno et al. (2011)) would have a prominent peak like structure in these phase bins. If we refer to the decomposed pulse profiles in Fig. 8 of Sasaki et al. (2012) for the decay phase of the outburst, we see that the two peaks in our overall pulse profile roughly correspond to the emission from the two poles. This lends support to the two pole possibility. However, we cannot rule out the fan beam / pencil beam as we have not examined the possible pulse profile and phase lags caused by such a possibility. Modelling these “wavy” phase lags, as attempted before (Ferrigno et al. 2011; Schönherr et al. 2014) could help resolve between these two possibilities. Pulse phase resolved spectra could have given additional indicators. However, the large time resolution of the Suzaku XIS data (of 2s) prevented us from getting the spectra at small phase bins of the pulsar spinning at 3.6s.

One way to confirm our hypothesis would be to use a single wide-band large area detector to make similar observations during the next outburst of this source. The constraint of using multiple instruments to bridge over the energy region having CRSFs and the region having the blackbody spectrum makes it very important to have these instruments cross-calibrated properly. While the cross-normalisation constants that we use compare favourably with the calibration carried out by Tsujimoto et al. (2011) for all instruments except Suzaku PIN-XIS, we do note that uncertainties in instrument cross-normalisations lead to uncertainties in the computed flux values and line equivalent widths estimated. Having fewer number of instruments to cover the range would then be ideal. For example, five to six snapshot observations over different luminosities using the XMM Newton and NuSTAR telescopes would definitely give a higher signal to noise data and lesser uncertainty in order to confirm or reject our hypothesis. Another way to do this would be the possible construction of a polarization spectrum. It is a well known fact that cyclotron resonant scattering has highly enhanced cross-sections for incident light polarized in the direction parallel to the local magnetic field, versus its cross-section for light polarized perpendicular to it. This enhancement occurs at the resonant scattering energy and drops off at other energies (see Becker & Wolff 2003 and references there-in). So, potentially a measurement of polarization in small energy bands can pin-point the energy range over which resonant scattering occurs, and effectively delineate the presence of actual CRSFs from those resulting due to incorrect spectral modelling. Future observations by detectors proposed for measuring the polarization in different energy bands (Paul et al. 2011; Havashida et al. 2014) will definitely help improve our understanding of this problem.

4.3 Conclusion

In this paper, we demonstrate the utility of having wide-band high signal to noise data by making use of Suzaku and RXTE satellite data. Using this, we point out the presence of a blackbody component. We also note the possible indications of two sets of CRSFs in this source. We note that having two sets of lines at 11 keV and 15 keV explains the reason for the observed anti-correlation in the fundamental CRSF energy with source luminosity. It additionally explains the reason for this source being the only known accretion powered pulsar to show 5 harmonics of the fundamental cyclotron line. Data from large area wide-band telescopes like NuSTAR and XMM or from telescopes which can give a polarization spectrum can potentially help solve this long open problem.
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Figure 8. Phase lag spectrum and the pulse profile. Bottom panel shows the pulse profile with shaded regions depicting the phase-bins over which correlation was computed. Top left panel is the phase-lag spectrum (and the correlation coefficient) for the Main peak (shaded dark gray in the pulse profile). Top right panel is the phase-lag spectrum from the secondary peak (shaded light gray in the pulse profile).

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APPENDIX A: FITTING THE CONTINUUM BLACKBODY

The blackbody component that we use for the continuum modelling does not consider effects of compton scattering of this component from the accreting plasma. [Farinelli et al. 2012] introduced an XSPEC model for implementing this. The results of using this comptonized blackbody are detailed in this section.

There are nine parameters required to describe the comptonized blackbody against two required for the simple blackbody. For getting the model to fit, we had to freeze the plasma parameters to the values obtained by previous attempts to model this source using such a bulk comptonization of seed photons. [This model fits for the plasma electron temperature (kT_e = 1.3 keV) and optical depth (τ = 0.41), velocity profile of the accreting plasma (two parameters, η = 0.5, β = 0.22), radius of accretion column (r_0 = 0.1) and albedo percentage (A = 1) from the NS surface (see Table 1 of Farinelli et al., 2012). The remaining free parameters are the blackbody temperature and the model normalization. The results of such a fit are summarized in Table A1. As seen on comparing these results with Table 3, we get blackbody seed temperatures to be similar to the model with a simple blackbody. However, the radius of the

6 http://adsabs.harvard.edu/abs/2013A%26A...553A.103F
component is indeed required to model the spectrum. Larger errors. This gives a strong indication that a blackbody drop to more reasonable values, although with smaller calculated from the normalization parameter. A blackbody component is indeed required to model the spectrum.

\[ a \]

Table A1. Blackbody temperature and radius for wide-band observations.

| Instrument        | Day (MJD) | \( kT \) (keV) | Radius (km) |
|-------------------|-----------|----------------|-------------|
| Swift / INTEGRAL  | 55736.34  | \( 1.19^{+0.26}_{-0.12} \) | \( 6.67^{+2.96}_{-4.32} \) |
| Suzaku / RXTE     | 55750.82  | \( 0.76^{+0.01}_{-0.01} \) | \( 6.21^{+1.15}_{-1.59} \) |

\( a \) calculated from the normalization parameter.

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