IRAC EXCESS IN DISTANT STAR-FORMING GALAXIES: TENTATIVE EVIDENCE FOR THE 3.3 µm POLYCYCLIC AROMATIC HYDROCARBON FEATURE?

B. Magnelli,1,2 R. R. Chary,2 A. Pope,3,4 D. Elbaz,4 G. Morrison,5,6 and M. Dickinson7

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ABSTRACT

We present evidence for the existence of an IRAC excess in the spectral energy distribution (SED) of five galaxies at 0.6 < z < 0.9 and one galaxy at z = 1.7. These six galaxies, located in the Great Observatories Origins Deep Survey field (GOODS-N), are star-forming since they present strong 6.2, 7.7, and, 11.3 µm polycylic aromatic hydrocarbon (PAH) lines in their Spitzer IRS mid-infrared spectra. We use a library of templates computed with PEGASE.2 to fit their multiwavelength photometry and derive their stellar continuum. Subtraction of the stellar continuum enables us to detect in five galaxies a significant excess in the IRAC band pass where the 3.3 µm PAH is expected (i.e., IRAC 5.8 µm for the range of redshifts considered here). We then assess if the physical origin of the IRAC excess is due to an obscured active galactic nucleus (AGN) or warm dust emission. For one galaxy evidence of an obscured AGN is found, while the remaining four do not exhibit any significant AGN activity. Possible contamination by warm dust continuum of unknown origin as found in the Galactic diffuse emission is discussed. The properties of such a continuum would have to be different from the local universe to explain the measured IRAC excess, but we cannot definitively rule out this possibility until its origin is understood. Assuming that the IRAC excess is dominated by the 3.3 µm PAH feature, we find good agreement with the observed 11.3 µm PAH line flux arising from the same C–H bending and stretching modes, consistent with model expectations. Finally, the IRAC excess appears to be correlated with the star formation rate in the galaxies. Hence it could provide a powerful diagnostic for measuring dusty star formation in z > 3 galaxies once the mid-infrared spectroscopic capabilities of the James Webb Space Telescope become available.

Subject headings: galaxies: evolution — galaxies: starburst — infrared: galaxies

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1. INTRODUCTION

Measuring the star formation history of galaxies as a function of redshift enables the build up of stellar mass in the universe to be constrained. Although there are systematic uncertainties between different star formation tracers, results from different studies (UV, IR) seem to converge on a flat or a gradual decrease in the star formation rate (SFR) between z = 3 and z = 1 followed by a steeper decline between z = 1 and z = 0 (e.g., Schiminovich et al. 2005; Chary & Elbaz 2001).

At high redshift (z > 3), evolution of the SFR still remains uncertain, primarily due to the poorly determined dust extinction correction. At these redshifts the SFR in galaxies is estimated using the UV luminosity which is strongly affected by dust extinction (Steidel et al. 1999). The extinction correction can be calculated using the UV slope (Meurer et al. 1999; Adelberger & Steidel 2000), but this technique is affected by several limitations. The UV slope can be influenced by the presence of an evolved stellar population and therefore overestimates the extinction correction. Moreover, it is well known that local galaxies harboring strong dusty star formation are opaque to UV radiation (e.g., Buat et al. 2005). This implies that the UV luminosity may not be a reliable tracer of the SFR in certain galaxies and that mid- and far-infrared tracers which are correlated with dust emission are required. The strongest dust emission features, such as those arising between 6 and 12 µm in polycyclic aromatic hydrocarbon (PAH) molecules, are redshifted out of the mid-infrared passband at z > 3. As a result, even deep 24 µm observations are insensitive to dusty star formation at these redshifts. Hence, a complete understanding of dust correction will require deep observations in the far-infrared/submillimeter regime or require accurate calibration of near-infrared (NIR) tracers of dusty star formation.

The James Webb Space Telescope (JWST) will obtain spectra between 5 and 27 µm and thereby detect the redshifted 3.3 µm PAH feature from galaxies at z > 3. The success of previous studies using monochromatic PAH luminosities as dusty SFR tracers (e.g., Chary & Elbaz 2001; Takeuchi et al. 2005; Brandl et al. 2006) motivates the calibration of this PAH emission feature as a SFR indicator.

Some studies have already investigated the possibility of using the 3.3 µm PAH feature as a SFR tracer and have revealed the presence of a correlation between L3.3 µm and the SFR in local galaxies (Mouri et al. 1990; Imanishi 2002; Imanishi et al. 2006). However, the existence of this 3.3 µm PAH feature at high redshifts needs to be confirmed and its line strength calibrated to the true SFR of galaxies.

In this paper we analyze the multiwavelength properties of six galaxies at z > 0.5 and assess the evidence for the presence of the 3.3 µm PAH feature. The galaxies display strong 6.2, 7.7, and 11.3 µm PAH emission in their Spitzer mid-infrared spectra and are hence expected to also show 3.3 µm PAH emission.
The sample of galaxies, situated in the Great Observatories Origin Deep Survey–North (GOODS-N; Dickinson et al. 2003a) field, has been observed in the optical by the Hubble Space Telescope (HST) and in the infrared by the Spitzer Space Telescope. The multiwavelength photometry is fitted with spectral energy distributions (SEDs) from the stellar population synthesis model PEGASE.2 (Fioc & Rocca-Volmerange 1997). The residual emission in the IRAC passband, where the 3.3 μm PAH signature should be present, is analyzed to determine its origin. We first investigate the possibility that this excess originates from an obscured active galactic nucleus (AGN) using the SED of NGC 1068. Then we assess the possibility that this excess originates from warm dust emission, i.e., the 3.3 μm PAH line emission and/or a warm dust continuum. To test the hypothesis that the IRAC excess is dominated by the 3.3 μm PAH line flux we compare the derived line flux to the 11.3 μm PAH flux measured in the IRS spectra. Since models predict that the 3.3 and 11.3 μm PAH lines originate from C−H modes (Li & Draine 2001), their line fluxes should be correlated.

The layout of the paper is as follows: The sample is presented in § 2, and the SED fits with PEGASE.2 and the determination of the IRAC excess is presented in § 3. Section 4 discusses the origin of the IRAC excess which at these wavelengths can be due to hot dust emission from an obscured AGN (§ 4.1), free-free and/or recombination line emission (§ 4.2) or, finally, to the 3.3 μm PAH broad emission line and/or a warm dust continuum (§§ 4.3 and 4.4). Our conclusions are summarized in § 5.

Throughout this paper we will use a cosmology with $H_0 = 71$ km s$^{-1}$ Mpc$^{-1}$, $\Omega_m = 0.73$, and $\Omega_\Lambda = 0.27$.

2. SAMPLE SELECTION

The 3.3 and 11.3 μm PAH lines arise from stretching and out-of-plane bending modes of the C−H bond, respectively (Duley & Williams 1981; Leger & Puget 1984; Li & Draine 2001). In order to obtain the best constraints on the 3.3 μm PAH emission we consider bright 24 μm sources in the GOODS-N field which have spectral coverage of the 11.3 μm PAH feature. We selected a sample of 21 galaxies with mid-infrared spectra from Infrared Spectrograph (IRS; Houck et al. 2004) observations in Spitzer programs GO2-20456 (PI: Chary) and 262 (PI: Helou). Data reduction includes cleaning rogue pixels, removing latent charge build-up, removing the sky and averaging the 2D files together (Pope et al. 2008). Spectral extraction was performed using a 2 pixel window in SPICE, and data were calibrated using the same extraction window on a standard star spectrum. For more details on the observations and data reduction for the data from GO2-20456, see Pope et al. (2008).

The data from PID 262 were taken in spectral mapping mode. As for the GO2-20456 data, the latent charge build up was fit and rogue pixels cleaned. A spectral cube was generated using the CUBISM tool (e.g., Smith et al. 2007). We ensured that we did not include the slit loss correction factor (SLCF) while making the cube because the sources in the field of view are point sources while the SLCF is used to correct for light diffracted into the slit due to a source which is extended wider than the spectrograph slit. The spectral cube was generated on a 2.5″ spatial grid. The spectrum for each source was extracted in a 3″$^2$ × 3″ aperture around the source. The flux density in the spectrum was shown in Table 1. The optical photometry are MAG_AUTO values from the public GOODS-N catalogs. The infrared photometry is measured in 4″ diameter apertures, with appropriate aperture corrections for the wings of the point-spread function. Photometric uncertainties are negligibly small (except for MIPS 3419, since the galaxies are very bright, and are dominated by a 5% systematic calibration uncertainty (Sirianni et al. 2005; Reach et al. 2005).

To determine the total infrared luminosity of these galaxies, we include the photometric constraints available from Spitzer 16, 24, and 70 μm observations from Teplitz et al. (2006), R. Chary et al. (in preparation), and Frayer et al. (2006), respectively. In order to test the presence of an AGN, we also consider Very Large Array (VLA) 1.4 GHz observations (G. Morrison et al. in preparation), and Chandra X-ray observations (Alexander et al. 2003).8

All galaxies have measured spectroscopic redshifts. For MIPS 4, MIPS 6, and MIPS 5581, redshifts are taken from the Team Keck Treasury Redshift Survey (TKRS; Wirth et al. 2004), for MIPS 5 and MIPS 7 redshifts are from Cowie et al. (2004), while for MIPS 3419 we derived a spectroscopic redshift from its IRS spectrum using the 9.7 μm silicate absorption feature. Five galaxies (the ones with ground-based spectroscopic redshift, i.e., MIPS 4, 5, 6, 7, and 5581) are at 0.64 < $z < 0.84$, while the sixth, with the IRS redshift (i.e., MIPS 3419), is at $z = 1.70$. The redshift for MIPS 3419 was measured by aligning a local galaxy template (Mrk 231) whose mid-IR spectral shape is similar to the extracted spectrum and redshifting it until it matches the observed spectrum. Since the 9.7 μm feature is quite shallow in MIPS 3419, the range of plausible redshifts in agreement with the data is 1.6 < $z < 1.9$. This is consistent with a visual analysis of the extracted spectrum which shows the flux density dropping to zero at ~25 μm corresponding to the 9.7 μm silicate feature at $z = 1.6$. The galaxies studied in this paper are star-forming galaxies since their Spitzer mid-infrared spectra display strong 7.7 and 11.3 μm PAH emission. Four out of six sources which have adequate wavelength coverage in their mid-infrared spectrum even display strong 6.2 μm emission. We fit their IRS spectra together with their broadband emission at 16, 24, and 70 μm, and 1.4 GHz (see

8 For a detailed description of all the ancillary data existing in the GOODS fields, we refer to the GOODS public World Wide Web site (http://www.stsci.edu/science/goods).
Table 2) with the SED library of Chary & Elbaz (2001) (see Pope et al. 2008 for details). We note that no aperture corrections are necessary since both the spectra and broadband observations integrate the entire galaxy. The inferred bolometric luminosities are shown in Table 2 and indicate that they are luminous infrared galaxies (LIRGs), i.e., with $L_{IR}(8–1000 \mu m) > 10^{11} L_{\odot}$ and SFR $\geq 17 M_{\odot}$ yr$^{-1}$ (using the SFR-$L_{IR}$ conversion law).

3. DATA ANALYSIS

We estimate the stellar continuum in the IRAC passbands by fitting stellar population synthesis models to the multiband photometry. The library of stellar emission used as templates is computed with PEGASE.2 (Fioc & Rocca-Volmerange 1997). Using the range of star formation, infall, and wind histories described...
in Table 3, and available on the Le PHARE World Wide Web site (S. Arnouts et al., in preparation), we construct an atlas of templates spanning galaxies of all Hubble types.

For each object, original PEGASE.2 templates are shifted to the spectroscopic redshift of the source and convolved with the transmission of each observed filter. We apply an age constraint to ensure that galaxies do not correspond to a template with an age greater than the age of the universe at its redshift. Finally we include extinction as a free parameter following a Calzetti law (Calzetti et al. 1994).

The templates are fit to all the BViz photometry and some but not all of the IRAC passbands. The fit of the stellar continuum excludes the IRAC bands where the 3.6 μm PAH signature may be present and also excludes longer wavelengths which are influenced by the rising dust continuum. Hence, for galaxies situated at 0.5 < z < 1, the IRAC 5.8 and 8.0 μm photometry is excluded from the fit (i.e., MIPS 4, MIPS 5, MIPS 6, MIPS 7, MIPS 5581) and for the galaxy situated at z = 1.7 (i.e., MIPS 3419), only the IRAC 8.0 μm photometry is excluded from the fit.

Applying all these constraints, we calculate the best fit using a $\chi^2$ minimization technique. For each object the best fit yields the stellar continuum corresponding to a specific combination of four parameters: type of galaxy (E, Sa, Sb, ...), age, normalization (i.e., stellar mass), and extinction (see Table 4). The result of the individual fitting is presented in Figure 2. For all objects the residual difference between the observed photometry and the best fit is below 15%.

Once the best fit is obtained a residual IRAC flux can be derived by computing the difference between the IRAC photometry and the stellar continuum (see Table 5). The resulting IRAC excess is dependent on the stellar continuum used to fit the data and thus relies on the specific value of our four parameters (type of galaxy, ages, mass, extinction). To assess this dependence and to estimate the error associated to each IRAC excess that we measure, we have undertaken a Monte Carlo approach. For each source, we randomly vary the photometry in each observed band by up to ±3 $\sigma_{\text{band}}$ where $\sigma_{\text{band}}$ is the photometric uncertainty in that band. Hence these new values for the photometry are still consistent with the observations. We recompute the best fit SED to these new data values and derive the inferred IRAC excess. For each object, this procedure is repeated 100 times. The error is then defined as the standard deviation of these 100 values. The result of the error estimation is presented in Table 5. These error bars account for both the photometric error bars and the range of stellar continuum models which fit the observed data.

We note that the Maraston (2005) models for stellar emission cannot be used for this analysis because the empirical spectra of thermally pulsing AGB stars, which are a key feature of the models, do not extend beyond rest-frame 2.5 μm. Usage of these models would result in an inferred line flux systematically higher by a factor ~2 than the one derived with PEGASE.2, due to this break at 2.5 μm.

The results, shown in Table 5, indicate that for all but one galaxy (MIPS 3419), we have detected a significant excess in the IRAC 5.8 μm passband. We note that the absence of an IRAC excess for MIPS 3419 could be due to a wrong redshift determination. Indeed, as discussed in § 2, the redshift of this source was only derived from its IRS spectrum. In the next section, we discuss the physical origin of the IRAC excess for the remaining five galaxies.

4. DISCUSSION ON THE ORIGIN OF THE IRAC EXCESS

The IRAC excess can originate from four different components: (1) hot dust emission from an obscured AGN, (2) free-free and hydrogen recombination line emission from ionized gas, (3) a possible continuum observed in the diffuse medium of our

| ID Object | $S_{16\,\mu m}$ ($\mu$Jy) | $S_{24\,\mu m}$ ($\mu$Jy) | $S_{70\,\mu m}$ ($\mu$Jy) | $S_{1.4\,GHz}$ ($\mu$Jy) | $S_{1.4\,GHz}^{\text{predicted}}$ ($\mu$Jy) | $L_{1.4\,GHz}^\text{rest frame}$ ($10^{23}$ W Hz$^{-1}$) | $\log (L_{IR})$ ($L_\odot$) |
|-----------|-----------------|-----------------|-----------------|-----------------|-----------------|-----------------|-----------------|
| MIPS 4……. | 777 ± 20 | 1221 ± 12 | 11.0 ± 0.66 | 161 ± 12 | 126 | 2.4 | 11.90 ± 0.05 |
| MIPS 5……. | 575 ± 11 | 750 ± 7 | 5.6 ± 0.67 | 91 ± 14 | 79 | 1.4 | 11.75 ± 0.04 |
| MIPS 6…….. | 398 ± 5 | 721 ± 7 | 11.0 ± 0.66 | 64 ± 16 | 75 | 1.0 | 11.70 ± 0.04 |
| MIPS 7…….. | 582 ± 10 | 832 ± 8 | 14.0 ± 0.70 | 104 ± 11 | 81 | 2.7 | 11.89 ± 0.05 |
| MIPS 3419….. | 54 ± 7 | 113 ± 6 | <3.0 | <25 | 19 | <4.0 | 11.76 ± 0.10 |
| MIPS 5581….. | 194 ± 5 | 201 ± 6 | <3.0 | 16 ± 5 | 17 | 0.6 | 11.22 ± 0.09 |

### Table 2
Infrared and Radio Flux Densities of the Sample

| Object ID | Type | Age (Gyr) | Stellar Mass ($M_\odot$) | $E(B - V)$ | $\chi^2$ |
|-----------|------|-----------|-----------------|----------|--------|
| MIPS 4……. | Sbc | 7 | $1.8 \times 10^{11}$ | 0.18 | 5.15 |
| MIPS 5……. | Sa  | 6 | $2.8 \times 10^{10}$ | 0.72 | 9.75 |
| MIPS 6…….. | Sbc | 5 | $7.3 \times 10^{10}$ | 0.03 | 8.68 |
| MIPS 7…….. | E   | 3 | $2.3 \times 10^{11}$ | 0.88 | 15.75 |
| MIPS 3419….. | Sd  | 3 | $2.6 \times 10^{10}$ | 2.73 | 6.48 |
| MIPS 5581….. | Sb  | 3 | $3.4 \times 10^{10}$ | 0.03 | 22.26 |

Notes.—PEGASE.2 scenarios are used as template parameters. SFR = $\nu^{-1}M_{\odot}$, and gas infall is simulated as $f(t) = \exp(-t/T_c)$; $\nu$ is effectively the ratio between the star formation timescale and the star formation efficiency. The initial mass function used in our scenarios is taken from Rana & Basu (1992).
Fig. 2.—Determination of the stellar continuum for six galaxies in our sample. Optical and NIR photometry used to fit the stellar continuum are shown by red squares, NIR photometry excluded from the fit are shown by red stars, and the IRS spectra by a continuous line (all IRS spectra have been smoothed to the same spectral resolution of $R \sim 45$ at $15 \, \mu$m, using a Gaussian kernel). The stellar continuum computed by PEGASE.2 is shown with the red dotted line. The convolution of the continuum with each passband is shown as the open circle (only wavelengths blueward of the shaded region have been used to determine the fits to the stellar continuum). The dashed vertical line represents the observed wavelength at which the redshifted 3.3 $\mu$m PAH features would be present and the shaded area represents the bandpass of the IRAC filter in which this PAH features falls. Thin vertical lines show the location of the 6.2, 7.7, and 11.3 $\mu$m PAH lines. The bottom of each panel shows the ratio of the residual error, defined as the difference between the observations and the final fit, and the photometric uncertainty in each passband. Errors obtained for the passbands used to fit the stellar continuum are shown as black squares while black stars represent errors obtained for the passbands excluded from the fit. The difference at red wavelengths is most likely due to the 3.3 $\mu$m PAH and the increasing contribution from the VSG continuum. [See the electronic edition of the Journal for a color version of this figure.]
Galaxy (Flagey et al. 2006; Lu 2004) and in some local star-forming galaxies (Lu et al. 2003), and finally (4) the 3.3 μm PAH feature itself. In the following we discuss the contribution of each of these components to the measured IRAC excess.

4.1. Obscured AGNs

A possible origin for the IRAC excess found in five galaxies of our sample could be hot dust emission from an obscured AGN. We assess this possibility by studying the X-ray, radio, and optical properties of these galaxies.

Using Chandra observations in GOODS-N (Alexander et al. 2003), we find that three galaxies of our sample (MIPS 4, 6, and 7) are detected in the soft X-rays (0.5–2 keV) and only one in the hard band (2–8 keV) (MIPS 6; see Table 6). The photon index $\Gamma = 1.06$ of MIPS 6 indicates the presence of an obscured AGN. Moreover, the fact that the resolved fraction of the X-ray background (CXB; Worsley et al. 2004) decreases with increasing ground (CXB; Worsley et al. 2004) decreases with increasing redshift. We observe this trend, which is consistent with observations and thus suggests that hot dust from an obscured AGN could explain the IRAC excess. One should note that the IRAC 8 μm line fluxes have been calculated with the original PEGASE.2 fit and errors have been calculated using the Monte Carlo approach. The SFR is derived from $L_{IR}$ using eq. (1).

Assuming that the excess observed in the IRAC passband is only due to hot dust from an obscured AGN, we calculate the normalization factor for the NGC 1068 SED corresponding to this excess. Then subtracting the AGN contribution from the multiwavelength photometry and using the fitting procedure described in §3, we fit the revised photometry with the PEGASE.2 population synthesis model. The result of these fits are presented in Figure 3.

For three objects, MIPS 4, 5, and 5581, these fits yield a predicted flux in the IRAC 8 μm passband which exceeds the observed photometry by a factor 6.2, 4.7, and 18.6 $\sigma$, respectively. These overestimations of the IRAC 8 μm photometry prove that MIPS 4, 5, and 5581 cannot harbor such an obscured AGN since their IRAC color (IRAC5.8/IRAC8.0) do not follow the typical colors of hot dust emission in NGC 1068. Moreover, to match the IRAC 8 μm photometry the obscured AGN, if present in these galaxies, would have to be at least a factor of 1.3, 1.25, and 2.0 less luminous, respectively. Then these obscured AGNs would not be luminous enough to explain the 5.8 μm IRAC excess. On the contrary, for MIPS 6 and 7, the renormalized NGC 1068 SED results in an IRAC 8 μm flux, which is consistent with observations and thus suggests that hot dust from an obscured AGN could explain the IRAC excess. One should note that these conclusions are dependent on the NIR spectral index of NGC 1068 (i.e., $\alpha = 2.25$). To assess this dependence we also attempted to fit the IRAC excess using the SED of Mrk 231 (Armus et al. 2007), which has a low spectral index of $\alpha = 0.25$.

### Table 5

**Inferred PAH Properties**

| Id Object | IRAC Excess ($\times 10^{-20}$ ergs s$^{-1}$ cm$^{-2}$ Hz$^{-1}$) | 3.3 μm Fine Flux ($\times 10^{-22}$ W cm$^{-2}$) | 11.3 μm Line Flux ($\times 10^{-22}$ W cm$^{-2}$) | EW$_{\text{rest frame}}$ (nm) | SFR (M$_{\odot}$ yr$^{-1}$) | $L_{3.3 \mu m}/L_{IR}$ ($\times 10^{-3}$) |
|-----------|-------------------------------------------------|------------------|-----------------|-----------------|-----------------|------------------|
| MIPS 4...| $3.7 \pm 0.7$ | $4.92 \pm 1.45$ | $4.67 \pm 0.58$ | 227 | 137.9 | $2.7 \pm 0.8$ |
| MIPS 5...| $3.2 \pm 0.8$ | $3.66 \pm 0.75$ | $5.87 \pm 0.68$ | 58 | 97.8 | $2.81 \pm 0.8$ |
| MIPS 6...| $2.0 \pm 0.3$ | $2.75 \pm 0.30$ | $3.69 \pm 0.69$ | 119 | 86.3 | $2.48 \pm 0.5$ |
| MIPS 7...| $1.4 \pm 0.5$ | $1.34 \pm 0.62$ | $4.16 \pm 1.66$ | 45 | 136.0 | $1.29 \pm 0.6$ |
| MIPS 3419| $0.0 \pm 0.1$ | $0.00 \pm 0.11$ | $\leq 0.18$ | $< 11$ | 99.5 | $< 1.9$ |
| MIPS 5581| $0.6 \pm 0.1$ | $0.67 \pm 0.15$ | $0.17 \pm 0.9$ | 106 | 28.5 | $3.5 \pm 1.0$ |

*Notes.—3.3 μm line fluxes have been calculated with the original PEGASE.2 fit and errors have been calculated using the Monte Carlo approach. The SFR is derived from $L_{IR}$ using eq. (1).*

### Table 6

**X-Ray Properties of the Sample**

| Object ID | Soft X-Ray (0.5–2 keV) | Hard X-Ray (2–8 keV) | Photon Index ($\Gamma$) |
|-----------|-----------------------|----------------------|------------------------|
|           | Observed* | Model*  | Observed* | Model*  |           |
| 4...      | 1.42     | 1.86    | $< 3.21$ | 2.33    | 1.45      |
| 5...      | $< 0.76$ | 1.31    | $< 7.49$ | 1.68    | $\ldots$ |
| 6...      | 0.65     | 1.22    | 2.54     | 1.48    | 1.06      |
| 7...      | 1.06     | 1.07    | $< 2.73$ | 1.32    | $> 1.35$  |
| 5581...   | $< 0.29$ | 0.24    | $< 1.55$ | 0.33    | $\ldots$ |

* Values are $\times 10^{-16}$ ergs s$^{-1}$ cm$^{-2}$. 

The X-ray properties of the sample are summarized in Table 6. The table shows the observed and model X-ray fluxes, divided into soft (0.5–2 keV) and hard (2–8 keV) bands. The photon index is also provided for each object.
Fig. 3.—Determination of the stellar continuum for five galaxies from which we have removed an AGN contribution proportional to the IRAC excess calculated in § 3. Lines and symbols are the same as in Fig. 2. The stellar continuum computed by PEGASE.2 and the AGN contribution are shown as dotted and dashed lines, respectively (as before, only wavelengths blueward of the shaded region have been used to determine the stellar continuum). The final fit, the sum of the stellar and the AGN contributions, convolved through each band pass, is shown as an open circle. The X-ray contribution of the normalized AGN is shown in each panel in erg s$^{-1}$ cm$^{-2}$. For sources 4, 5, and 5581, the IRAC excess is inconsistent with being due to an obscured AGN (see text for details). [See the electronic edition of the Journal for a color version of this figure.]
We find that the AGN fit exceeds the observed 8 µm IRAC flux for MIPS 4, 5, and 5581 by a value of 4.1, 3.0, and 9.9 σ, respectively. Although these values are a factor of 2 smaller than the ones derived for NGC 1068 they also weaken the possibility that hot dust emission can explain the IRAC excess.

We now evaluate if the observed X-ray emission from these sources is consistent with this simple model. We assume that the X-ray emission is the sum of starburst and AGN contributions, and compare with the Chandra observations (see Table 6). Since our galaxies show strong PAH features, the bolometric luminosity calculated in § 2, is dominated by star formation. We calculate the starburst X-ray contribution using the relationship between the SFR and L_{IR} (Kennicutt 1998) and the relationship between the SFR and L_{0.5-8 keV} (Bauer et al. 2002):

\[ \text{SFR} \left[ M_\odot \text{yr}^{-1} \right] = 4.5 \times 10^{-44} L_{IR} \left[ \text{ergs s}^{-1} \right], \]

(1)

\[ \text{SFR} \left[ M_\odot \text{yr}^{-1} \right] = 1.7 \times 10^{-43} L_{1.07}^{0.5-8 \text{keV}} \left[ \text{ergs s}^{-1} \right], \]

(2)

Using \( \Gamma = 1.9 \) as the typical starburst photon index, we estimate the soft and hard X-ray rest-frame luminosity (\( L_{\text{rest frame}} \)) predicted for the starburst. The observed flux in each band (\( f_{\text{observed}} / f_{\text{soft/hard}} \)) is then derived by applying a K-correction given by eq. (2) of Bauer et al. (2002):

\[ f_{\text{observed}} \left[ \text{ergs s}^{-1} \text{cm}^{-2} \right] = \frac{L_{\text{rest frame}}^{\text{soft/hard}} \left[ \text{ergs s}^{-1} \right]}{4\pi d_l^{-2} (1+z)^{1/2}}, \]

(3)

where \( d_l \) is the luminosity distance. For the AGN contribution, we adopt \( L_{2-10 \text{keV}} = 2.8 \times 10^{44} \text{ergs cm}^{-2} \text{s}^{-1} \) and \( \Gamma = 1.2 \) (Ogle et al. 2003) for NGC 1068. From this, we predict the emission of our objects in the soft and hard X-ray bands (cf. Fig. 3) using the NGC 1068 normalization factor calculated from the IRAC excess and the K-correction given by equation (3).

For all galaxies the total (AGN+starburst contribution) soft X-ray emission, dominated by the starburst contribution (\( L_{\text{soft}} / L_{\text{AGN}} \sim 10 \)), is in agreement, within the error, with the observed soft X-ray flux by a factor of 1.3, 1.7, 1.8, and 1.0 for MIPS 4, 5, 6, and 7, respectively, and under the Chandra threshold for MIPS 5581 (see Table 6).

For MIPS 4, 5, 6, and 5581, the total hard X-ray prediction is under the Chandra threshold and thus in agreement with observations. For MIPS 6, our hard X-ray prediction (1.48 \times 10^{-16} \text{ergs s}^{-1} \text{cm}^{-2}) is a factor of 1.7 below the observation (2.54 \times 10^{-16} \text{ergs s}^{-1} \text{cm}^{-2}) but still, within the error, in agreement. This suggests that the obscured AGN harbored by this galaxy is significantly more luminous than for NGC 1068 in the X-rays.

In summary, the X-ray analysis reveals the presence of an obscured AGN in MIPS 6 but is not conclusive for MIPS 4, 5, 7, and 5581 since either the presence or absence of an obscured AGN would be in agreement with the data. The X-ray predictions using the Mrk 231 SED are similar to those derived using the SED of NGC 1068. Hence, our conclusions are independent of the AGN SED adopted to fit the data.

Using the Chary & Elbaz library which follows the local radio/far-infrared correlation we predict \( L_{1.4 \text{GHz}} \) using the 24 µm luminosity (see Table 2). For MIPS 4, 5, and 5581 we find \( L_{1.4 \text{GHz}} / L_{\text{IRAC 8}} \sim 0.8 \), and for MIPS 6 and MIPS 7, which may be powered by an AGN, we find \( L_{1.4 \text{GHz}} / L_{\text{IRAC 8}} \sim 1.2 \) and 0.8, respectively. This agreement between predictions and observations suggests that the infrared emission of all these galaxies is dominated by star formation. Finally, assuming a spectral index (\( \alpha \text{in F}_\nu \propto \nu^{-\alpha} \)) of 0.8 we calculate for each galaxy \( L_{1.4 \text{GHz}} \) (see Table 2). These values are then compared to 5 \times 10^{23} \text{W Hz}^{-1}, which is the typical radio luminosity used to distinguish between a starburst and an AGN population at \( z \sim 0.7 \) (Yun et al. 2001; Cowie et al. 2004). For each galaxy in our sample, \( L_{1.4 \text{GHz}} \) is below 5 \times 10^{23} \text{W Hz}^{-1}, which is consistent with the assumption that our galaxies are star-formation-dominated systems.

Finally, we analyze the optical spectra of three galaxies (MIPS 4, 6, and 5581) available in the TKRS database (Wirth et al. 2004). No clear evidence of high-ionization lines such as [Ne III] or [Ne V] or large [O III]/Hβ ratio are found, arguing against the possibility that the IRAC excess of these three galaxies is dominated from AGN activity. However, we note that this analysis includes MIPS 6 which shows clear evidence of an AGN in its X-ray emission.

The conclusion of the SED fitting, X-ray, optical, and radio diagnostics on the presence of an obscured AGN in these galaxies can be summarized as follows:

1. For MIPS 4, 5, and 5581 the disagreement between the IRAC 8 µm expected for a prototypical obscured AGN and the observations suggests that these galaxies do not harbor an AGN. Moreover, for MIPS 4 and 5581, this conclusion is also supported by the agreement found between their radio and mid-infrared luminosities as well as the absence of AGN signature in their optical spectra.

2. For MIPS 7, the IRAC excess could be reproduced by a buried AGN, but we found no other evidence for an AGN from either its optical or its X-ray properties.

3. Finally, MIPS 6 presents marginal evidence for the presence of an obscured AGN in its X-ray emission (1.7 times larger than typically expected for star formation). Its radio luminosity and optical spectrum are consistent with a star-formation-dominated system. In the following, we will discuss the case of this galaxy separately since we cannot rule out its contamination by an obscured AGN.

### 4.2. Free-Free and Gas Lines

The total contribution of free-free and recombination line emission have been estimated by Flagey et al. (2006) to be about 1% and 3% in the IRAC 3.6 µm (i.e., the residual IRAC 5.8 µm emission for MIPS 4, 5, 6, 7, and 5581 or the residual IRAC 8.0 µm for MIPS 3419) and in the IRAC 4.8 µm (i.e., the residual IRAC 8.0 µm for MIPS 4, 5, 6, 7, and 5581) channels, respectively. Even in the extreme case where these contributions would reach 3% and 11% (Flagey et al. 2006), the free-free and line emission contribution to the IRAC excess would be negligible. Furthermore if the IRAC excess was due to free-free emission, it would be accompanied by a comparably high value at shorter wavelengths (i.e., in the 4.5 µm IRAC channel) due to the blue spectrum of free-free emission. The flux density in the 3.6 and 4.5 µm IRAC channels in MIPS 4, 5, 6, 7, and 5581 are consistent with being dominated by stellar emission (see Fig. 2), excluding this possibility.

### 4.3. Warm Dust Continuum

A continuum underlying the 3.3 µm PAH feature has been detected in the Galactic diffuse medium (Flagey et al. 2006; Lu et al. 2004) and in some local star-forming galaxies (Lu et al. 2003). There, the NIR continuum, well fitted by a modified blackbody with temperature spanning the range 700–1500 K and a \( \lambda^{-2} \) emissivity law, contributes to about 70% of the IRAC 3.6 µm flux. They also conclude that the IRAC 4.5 µm channel is totally dominated by this NIR continuum.
Assuming that the IRAC excess, which we measure here in distant galaxies, is due to such a continuum, the $f^\text{predicted}_\nu/(8.0 \, \mu\text{m})$ for MIPS 4, 5, 6, 7, and 5581 would span the range [0.97–2.07]. We can assume these color predictions to be lower limits since they assume no contribution of the 3.3 $\mu$m PAH feature in the IRAC 5.8 $\mu$m passband. The observed ratios after subtracting the stellar contribution (i.e., $[f^\text{observed}_\nu(5.8 \, \mu\text{m}) - f^\text{PEGASE}_\nu(5.8 \, \mu\text{m})]/[f^\text{observed}_\nu(8.0 \, \mu\text{m}) - f^\text{PEGASE}_\nu(8.0 \, \mu\text{m})]$) are 0.64, 0.64, 0.45, 0.32, and 1.08 for MIPS 4, 5, 6, 7, and 5581, respectively. Although one galaxy is marginally consistent with this warm dust continuum (MIPS 5581), the remaining four correspond to a modified blackbody temperature of 550 K, and hence are colder than the temperature range of the NIR continuum found in Flagey et al. (2006). As a result, it is unlikely, unless the shape of the continuum were different from anything known, that the IRAC 5.8 $\mu$m passband is dominated by the same NIR continuum in MIPS 4, 5, 6, and 7.

The physical origin of such a high-temperature continuum remains unclear, and we cannot totally rule out that it reaches a different temperature range in distant highly star-forming galaxies. NIR spectroscopy (L-band, 3–4 $\mu$m) of a sample of 24 local actively star-forming galaxies (Imanishi & Dudley 2000) does exhibit some contribution from a hot continuum, which could, in principle, serve as a better reference for the present sample of distant star-forming galaxies. We find that their continuum can be reproduced by a colder modified blackbody than the one found in Flagey et al. (2006) since it reaches a temperature of $T \sim 400$ K. This colder continuum, which is more consistent with the IRAC temperature of the present sample, would contribute for 60% of the IRAC excess. However, this sample consists only of the nuclei of nearby ultraluminous infrared galaxies (ULIRGs) obtained using the Subaru IRCS NIR spectrograph. These galaxies were chosen as candidate AGNs, the goal of the authors being to look for buried AGNs, which they claim to find with various intensities in all objects. Hence, the warm dust continuum observed in the line sample can be highly contaminated by AGN activity.

4.4. 3.3 $\mu$m PAH Feature

Assuming that the IRAC excess is due to the 3.3 $\mu$m PAH feature, we show that the inferred line fluxes are consistent with expectations from dust models. To infer the 3.3 $\mu$m PAH line flux from the IRAC excess, we adopt a Gaussian profile for the PAH line. Then we convolve this profile with the IRAC filter curve to obtain the normalization factor for the Gaussian profile corresponding to the IRAC excess. Finally, the true flux of the 3.3 $\mu$m PAH line is obtained by integrating over the normalized Gaussian profile. The estimated value of the 3.3 $\mu$m PAH line flux calculated for each galaxy is shown in Table 5. We note that the line flux inferred by adopting a Drude profile is only 1.12 times the value derived for a Gaussian profile. Therefore, our conclusion does not depend on this choice. We also note that the extinction derived from the stellar population fits to the photometry (see Table 4) reaches an average value of $A(3.3 \, \mu\text{m}) \approx 0.14$. Hence, the extinction correction to the inferred 3.3 $\mu$m PAH line flux is smaller than the uncertainty in the flux and are therefore negligible.

Models predict that the 3.3 and 11.3 $\mu$m lines originate in the same C–H bond (Li & Draine 2001; Duley & Williams 1981), which results in a linear correlation between these two lines. Hence, finding a correlation between the inferred 3.3 $\mu$m line flux and 11.3 $\mu$m line flux, measured from the IRS spectra, would support the assumption of a negligible continuum.
Imanishi & Dudley (2000). This aperture correction relies on the assumption that the nucleus and the whole galaxy have the same spectrum, which may not be the case for AGN-dominated nuclei. In fact, we note that this aperture corrections would provide only a lower limit to \(L_{3.3 \mu m}\) due to the low 3.3 \(\mu m\) equivalent widths observed in obscured AGNs.

The values of \(L_{3.3 \mu m}\) inferred for five out of six galaxies (MIPS 4, 5, 6, 7, and 3419) are consistent with the one predicted by the Li & Draine (2001) model from the measured \(L_{11.3 \mu m}\). This agreement supports the hypothesis that the IRAC excess observed in these galaxies is due to the 3.3 \(\mu m\) PAH feature. This true also for MIPS 6 and MIPS 7, suggesting that their potential AGN contribution (see § 4.1) does not dominate the IRAC flux excess. For MIPS 5581, the inferred 3.3 \(\mu m\) PAH line is overestimated by a factor of \(1.5-4\) with respect to model expectations from the \(L_{11.3 \mu m}\). This can be explained by the presence of a significant NIR continuum as also revealed by the color diagnostic presented in § 4.3. The assumption that the NIR continuum is negligible compared to the line leads to an overestimation of the inferred 3.3 \(\mu m\) PAH line.

The presence of strong PAH features in the IRS spectra of each galaxy of the sample, except for MIPS 3419, indicates that they must be star-formation-dominated. The \(L_{3.3 \mu m}/L_{IR}\) ratio in local star-forming galaxies has been found to be \(\sim 10^{-3}\) (Mouri et al. 1990) with a factor of 2–3 scatter.

For MIPS 4, 5, 6, 7, and 5581 we find \(L_{3.3 \mu m}/L_{IR,3} \times \sim 10^{-3}\), which is, within the uncertainties of the relation, consistent with the local ratio. For MIPS 3419, the upper limit found for the 3.3 \(\mu m\) PAH line flux is still consistent with the local ratio of \(L_{3.3 \mu m}/L_{IR}\).

The agreement of the \(L_{3.3 \mu m}/L_{11.3 \mu m}\) ratio with standard dust models and of the \(L_{3.3 \mu m}/L_{IR}\) ratio with local observations favors the hypothesis that the IRAC excess is effectively due to the 3.3 \(\mu m\) PAH emission. We also note that even in the case where the IRAC excess would be due to the combined emission of a continuum and the 3.3 \(\mu m\) PAH, the study of nearby star-forming galaxies by Lu et al. (2003) showed that both emissions probably originate from the same carriers. Indeed, they find that the 7.7 \(\mu m\) PAH equivalent width does not vary in their sample of galaxies. This trend was not corroborated by Flagey et al. (2006), but the latter study remains limited to the diffuse ISM of the Milky Way. In light of Lu et al. (2003) results the information carried by the 3.3 \(\mu m\) PAH line flux alone or by the continuum and the 3.3 \(\mu m\) PAH complex would be equivalent.

5. CONCLUSION

We have investigated the presence of an excess in the observed IRAC flux densities through analysis of the multiwavelength photometry and spectroscopy of six galaxies in the redshift range 0.5 < z < 1.8. Using a fit to the optical and NIR photometry, we determined the stellar continuum of each source with PEGASE.2. The difference between the stellar continuum and the IRAC flux in the passband expected to harbor the 3.3 \(\mu m\) PAH feature is computed and analyzed. For five galaxies we found a significant IRAC excess. For the other galaxy, we were only able to determine an upper limit.

We investigated the possibility that the measured IRAC excess could be explained by the presence of an obscured AGN using the SED of NGC 1068. For 1 galaxy (MIPS 6) we find evidence for a possible contamination by an obscured AGN. For four galaxies (MIPS 4, 5, 7, and 5581), no evidence was found that the IRAC excess could result from the presence of an obscured AGN.

We have investigated the origin of the IRAC excess as potentially due to the presence of the 3.3 \(\mu m\) PAH feature and/or a warm dust continuum. Local observations (Flagey et al. 2006; Lu et al. 2003) find evidence for a warm dust continuum of unknown origin which could dominate the IRAC excess. The NIR colors of the IRAC excess cannot be reproduced by such a continuum unless one assumes a lower temperature. However, the 3.3 \(\mu m\) PAH line flux inferred directly from the IRAC excess is consistent with standard dust models. Indeed, we find that the inferred 3.3 \(\mu m\) PAH line flux is compatible with the 11.3 \(\mu m\) PAH line flux measured in the IRS spectrum. This suggests that the IRAC excess is effectively due to the 3.3 \(\mu m\) PAH feature.

However, we note that even if the IRAC excess would arise from the combination of the 3.3 \(\mu m\) PAH line flux and a warm dust continuum such as that found in local star-forming galaxies by Lu et al. (2003) both originate from the same carriers, as indicated by a constant PAH/continuum flux ratio. In that case, as well as in the case where the 3.3 \(\mu m\) PAH line flux alone dominates the IRAC excess, we conclude that this emission provides a powerful diagnostic for measuring dusty star formation rates in galaxies at z > 3 using the JWST. We find that \(\nu_{rest, 3.6 \mu m}/L_{IR}\) is nearby constant for all five galaxies, i.e., \((1.0 \pm 0.3) \times 10^{-3}\) with \(L_{IR}\) spanning the range \([1.5 \times 10^{11} \text{ to } 8 \times 10^{11}]\).

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REFERENCES

Adelberger, K. L., & Steidel, C. C. 2000, ApJ, 544, 218
Alexander, D. M., Bauer, et al. 2003, AJ, 125, 383
Armus, L., et al. 2007, ApJ, 656, 148
Barger, A. J., Cowie, L. L., & Wang, W.-H. 2007, ApJ, 654, 764
Bauer, F. E., Alexander, D. M., Brandt, W. N., Hornschemeier, A. E., Vignali, C., Garme, G. P., & Schneider, D. P. 2002, AJ, 124, 2351
Brandl, B. R., et al. 2006, ApJ, 653, 1129
Buat, V., et al. 2005, ApJ, 619, L51
Calzetti, D., Kinney, A. L., & Storchi-Bergmann, T. 1994, ApJ, 429, 582
Chary, R., & Elbaz, D. 2001, ApJ, 556, 562
Cowie, L. L., Barger, A. J., Hu, E. M., Capak, P., & Songaila, A. 2004, AJ, 127, 3137
Dickinson, M., Giavalisco, M., & The GOOGS Team. 2003a, in The Mass of Galaxies at Low and High Redshift, ed. R. Bender & A. Renzini (Berlin: Springer), 324
Draine, B. T., & Li, A. 2007, ApJ, 657, 810
Duley, W. W., & Williams, D. A. 1981, MNRAS, 196, 269
Fazio, G. G., et al. 2004, ApJS, 154, 10
Fioc, M., & Rocca-Volmerange, B. 1997, A&A, 326, 950
Flagey, N., Boulanger, F., Verstraete, L., Miville Descheˆnes, M. A., Noriega Crespo, A., & Reech, W. T. 2006, A&A, 453, 969
Frayer, D. T., et al. 2006, ApJ, 647, L9
Galliano, E., Alloin, D., Granato, G. L., & Villar-Martın, M. 2003, A&A, 412, 615

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