The GALAH survey: a catalog of carbon enhanced stars and CEMP candidates

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ABSTRACT
Swan bands – characteristic molecular absorption features of the C₂ molecule – are a spectroscopic signature of carbon enhanced stars. They can also be used to identify notorious carbon enhanced metal-poor (CEMP) stars. GALAH (GALactic Archaeology with Hermes) is a magnitude limited survey of stars producing high resolution, high signal-to-noise spectra. We use 627,708 GALAH spectra to search for carbon enhanced stars with a supervised and unsupervised classification algorithm, relying on the imprint of the Swan bands. We identified 918 carbon enhanced stars, including 12 already described in the literature. An unbiased selection function of the GALAH survey allows us to perform a population study of carbon enhanced stars. Most of them are giants, out of which we find 28 CEMP candidates. Large fraction of stars with repeated observations show variation in radial velocity, hinting that there is a large fraction of variables among carbon enhanced stars. 32 of detected stars also show strong Lithium enhancement in their spectra.

Key words: methods: data analysis – stars: carbon – stars: abundances – catalogues

1 INTRODUCTION
Chemically peculiar stars whose spectra are dominated by carbon molecular bands were first identified by Secchi (1869). Their spectra are characterized by enhanced carbon absorption bands of CH, CN, SiC₂, and C₂ molecules, also known as Swan bands. Possible sources of enhancement are dredge-up events in evolved stars (Iben 1983), enrichment by a carbon rich stellar wind originating in a pulsating asymptotic giant branch (AGB) star, which settles on a main sequence companion (Han et al. 1995) or the result of primordial enrichment. Historically, high latitude carbon stars,
presumed to be giants, were used as distant kinematic and dynamical probes to measure the Galactic rotation curve (Battinelli et al. 2013), velocity dispersions in the Galactic halo (Bothun et al. 1991), and to trace the gravitational potential of the Galaxy.

Because of their strong spectral features, the most prominent candidates can easily be identified from large photometric (Margon et al. 2002; Downes et al. 2004) surveys. Specific photometric systems (Griffin & Redman 1960; McClure & van den Bergh 1968; Haggkvist & Oja 1970) were defined in the past to discover and further classify stars with enhanced carbon features in their spectra. Specifics of those photometric systems were cataloged, compared and homogenized by Moro & Munari (2000) and Fiorucci & Munari (2003).

Other useful data come from low resolution spectroscopic surveys, whose exploration identified from few hundred to few thousand of those objects (Christlieb et al. 2001; Green 2013; Ji et al. 2016; Li et al. 2018). High resolution spectroscopy is required to search for candidates with less pronounced molecular absorption features or to determine the stellar chemical composition. However, the number of stars with detailed studies of chemical abundances, is low. Cohen et al. (2006) performed detailed abundance analysis for 16 carbon stars with extremely low metallicity.

Today, the most sought after, of all carbon enhanced stars, are carbon enhanced metal poor (CEMP) candidates whose fraction, among metal poor stars, increases with decreasing metallicity [M/H] (Lee et al. 2013). Amongst these, the ones lying around the main-sequence turn-off are expected to be of particular importance, as they may have accreted enough material from their AGB companion to produce an observable change in their atmospheric chemical composition. This could provide insight into the production efficiency of neutron capture elements in AGB stars (Aoki et al. 2007). The exact origin and underlying physical processes governing multiple classes seen in CEMP stars are not yet understood and are topics of ongoing research (Yoon et al. 2016; Cruz et al. 2018).

In this work, we propose a novel approach for the classification of carbon enhanced stars using high resolution stellar spectra covering parts of the visible domain. The goal is to identify a representative sample of carbon enhanced stars, which can be used as an input to population studies. The paper is organized as follows; we start with a brief discussion of our spectroscopic observations and their reduction (Section 2), which is followed by the description of the used algorithm for the detection of carbon enhanced stars in Section 3. Properties of the classified objects are investigated in Section 4, CEMP candidates are a focus of Section 5, with Section 6 describing a follow-up study for one of them. Final remarks are given in Section 7.

2 DATA

The analysed set of stellar spectra was acquired by the High Efficiency and Resolution Multi-Element Spectrograph (HERMES), a fibre-fed multi-object spectrograph on the 3.9 m Anglo-Australian Telescope (AAT) of the Australian Astronomical Observatory. HERMES (??) can simultaneously record spectra from up to 392 fibres distributed over 2° field of the night sky, with an additional 8 fibres used for the telescope guiding. The spectrograph has a resolving power of R ~ 28,000 and consists of four spectral arms centered at 4800 Å, 5761 Å, 6610 Å, and 7740 Å, together covering approximately 1000 Å, including the Hα and Hβ lines. Three dichroic beam splitters are used to separate incoming light into four separated colour beams that are analysed independently. The spectrograph can typically achieve a signal to noise ratio (SNR) ~100 per resolution element at magnitude V=14 in the red arm during a 1 hour long exposure.

Spectra used in this study have been taken from multiple different observing programs using this spectrograph: the GALactic Archaeology with HERMES (GALAH) pilot survey (Duong et al. 2018), the main GALAH survey (De Silva et al. 2015), the K2-HERMES survey (Wittenmyer et al. 2016), and the TESS-HERMES survey (Sharma et al. 2018). Most of those observing programs exclude fields close to the Galactic plane (due to problems with high stellar density and Galactic extinction) or far away from it (not enough suitable targets to use all fibres), employ subtle different selection functions (position, limiting magnitude, crowding requirement, photometric quality), but share the same observing procedures, reduction and analysis pipeline (internal version 5.3, Kos et al. 2017). All programs, except the pilot survey, are magnitude limited. This leads to an unbiased sample of stars distributed mostly across the southern sky that can be used for different population studies. Additionally, all objects from different observing programs are analysed with the same procedure named The Cannon (internal version 180325, Ness et al. 2015; Buder et al. 2018), so their stellar parameters are determined in a consistent manner and are hence comparable across the different programs. The Cannon algorithm employs a data driven interpolation approach trained on a set of high quality benchmark stars spanning the majority of the stellar parameters space (for details, see Buder et al. 2018).

The spectrum synthesis code Spectroscopy Made Easy (Valenti & Piskunov 1996; Piskunov & Valenti 2017) was run on the spectra in the training set to determine their stellar parameters and atmospheric chemical abundances. After training The Cannon model, it was run independently on every observed spectrum to determine its stellar parameters. Every determined parameter is accompanied by a quality flag identifying its usefulness and possible problems with grid interpolation. Therefore we can easily remove parameters that were determined for stellar spectra far away from the training set or with possible problems in any of the HERMES spectral arms. An interpolation method proved beneficial in determining stellar physical parameters and up to 23 abundances (extendable up to 31 in future versions) for every analysed star, as described in Buder et al. (2018). The parameter [Fe/H] determined by The Cannon refers to an iron abundance and not overall metallicity [M/H] as often used in the literature. As both notations are used interchangeably in the literature to describe stellar metallicity, care must be taken when comparing those two measurements.

The Cannon approach has an advantage of treating stars with a peculiar composition, such as carbon enhanced stars, in a consistent manner with common objects, thus avoiding arbitrary jumps or offsets that would depend on a degree of peculiarity or correctness of physics of the underlying model. As a data driven approach, it only projects
the learned training set onto the whole survey based on quadratic relations to stellar parameters.

Our dataset consists of 627,708 successfully reduced spectra of 576,229 stars observed between November 2013 and February 2018.

3 DETECTION PROCEDURE

To search for carbon enhanced stars in the GALAH dataset we focused on spectral features that can be clearly distinguished and are known markers of carbon enhancement. Instead of using one very weak atomic carbon absorption line (at 6587.61 Å), used by The Cannon to determine [C/Fe] abundance, we focused on a region between 4718 and 4760 Å observed in the blue arm, which covers 4718–4903 Å. In this range we can, depending on the radial velocity of the star, observe at least four Swan band features (John-son 1927) with their band heads located at approximately 4715 Å, 4722 Å, 4737 Å, and 4745 Å.

Carbon enhancement is observable in spectra as a strong additional absorption feature (Figure 1) that is the strongest at the wavelength of the band’s head. After that its power gradually decreases with decreasing wavelength. The most prominent and accessible for all of the spectra is a feature located at 4737 Å, produced by a 12C12C molecule. If other carbon features, like the one produced by a 13C12C molecule at 4745 Å (shown in Figure A3) are present in the spectrum, the carbon isotope ratio 12C/13C in a star can be determined. It’s determination was not attempted in the scope of this paper.

Detection of spectral features in was tackled using two different classification procedures. First a supervised procedure was used to identify the most prominent spectra with carbon enhancement. It is based on the assumption that we know where in the spectra those features are located and how they behave. This was augmented with an unsupervised dimensionality reduction algorithm that had no prior knowledge about the desired outcome. The goal of a dimensionality reduction was to transform n-dimensional spectra onto a 2D plane where differences between them are easier to analyse. The unsupervised algorithm was able to discern the majority of carbon enhanced spectra from the rest of the dataset and enabled us to discover spectra with less prominent carbon enhancement features.

3.1 Supervised classification

To search for any additional absorption features that are usually not found in spectra of chemically normal stars, we first built a spectral library of median spectra based on rough estimates of stellar physical parameters derived by the automatic reduction pipeline, described in detail by Kos et al. (2017). The median spectrum for every observed spectrum in our dataset was computed from physically similar spectra with stellar parameters in the range of $T_{\text{eff}}=\pm 75$ K, log $g=\pm 0.125$ dex and [Fe/H]=±0.05 dex around the stellar parameters of the investigated spectrum. The median spectrum was calculated only for observed targets with at least 5 similar spectra in the defined parameter range with minimal SNR=15 per resolution element as determined for the blue arm. All considered spectra were resampled to a common wavelength grid with 0.04 Å wide bins and then median combined. The normalization of the spectra along with the radial velocity determination and the corresponding shift to the rest frame was performed by the automatic reduction pipeline (Kos et al. 2017). No reduction step seems to be problematic for carbon enhanced stars. The normalization procedure is done using a polynomial of low order that is not affected by the Swan band features or similar structure. The radial velocity of a star is determined as an average of radial velocities that were independently determined for the blue, green and red spectral arm. If one of the arms has a radial velocity deviating for more than two times the difference between the other two, it is excluded from the average (further details in Kos et al. 2017). Therefore the final velocity should be correct even if one of those arms includes features that are not found in the set of reference spectra used in the cross-correlation.

With the limitation of at least 5 spectra used for the computation of the median spectrum, some possibly carbon enhanced stars are excluded from the supervised classification. The final number of spectra analysed by this method was 558,053.

Spectra, for which we were able to determine the median spectrum of physically similar objects, were analysed further. In the next step, we tried to determine possible carbon enhancement by calculating flux difference and flux division between the observed stellar and median spectra, as shown in Figure 1.

In order to describe the position, shape and amplitude of the Swan feature with its head at 4737 Å, we fitted the function that is based on a Log Gamma ($\log \Gamma$) distribution. The distribution, with three free parameters, was fitted to the division curve, where the Swan feature is most pronounced. Division curve, shown in the bottom panel of Figure 1, was computed by dividing observed spectrum with its corresponding median spectrum. The fitted function $f$ can be written as:

$$f(\lambda) = f_0 - \log \Gamma(\lambda, S, A_0, A).$$

The shape of the curve is defined by an offset $f_0$, shape parameter $S$, center wavelength $\lambda_0$ and amplitude $A$ of $\log \Gamma$ distribution, where $\lambda$ represents rest wavelengths of the observed spectrum. This function was selected because of its sharp rise followed by the gradual descent that matches well with the shape of a residual absorption observed in the Swan regions. Steepness of the rising part is determined by the parameter $S$ (lower value indicates steeper raise) and its vertical scaling by the parameter $A$. We are not aware of any other profile shapes used for fitting Swan bands in the literature.

To narrow down possible solutions for the best fitting curve, we used the following priors and limits. The initial value for the parameter $f_0$ was set to a median of all pixel values in division curve and allowed to vary between 0.5 and 1.5. The limiting values are however never reached. The center of the $\log \Gamma$ distribution $\lambda_0$ was set to 4737 Å and was allowed to vary by 2 Å. Wavelength limits were set to minimize the number of mis-detections as the fit could get stuck into the nearby spectral absorption lines not present in the median spectra or into spectral feature caused by problem in the spectral data reduction as shown by the example in...
Figure 1. Example of a metal poor carbon enhanced candidate with strong Swan absorption feature at 4737 Å, caused by the carbon C_2 molecules. The first panel shows the stellar spectrum in blue and a corresponding reference median spectrum in red. The reference median spectrum was computed as the per pixel median flux value of spectra with similar stellar parameters as the spectrum shown in blue. The spectral difference (second panel) and division (third panel) were computed from those spectra. The middle panel shows in orange a linear fit to the spectral difference that was used to identify spectra with the reduction problems on the blue border of the spectral range. The median value of the spectral difference is given by the black horizontal line. The orange curve in the last panel shows a fit that was used to determine the strength of observed carbon feature. The shown spectrum belongs to a star with a Two Micron All-Sky Survey (2MASS) identifier J11494247+0051023 and iron abundance [Fe/H] of $-1.17$ dex as determined by The Cannon.

Figure A5. We did not set any limits on parameters A and S in order to catch fitted solutions describing a curve different to the expected shape of the molecular absorption band.

By integrating the surface between the offset $f_0$ and the fitted curve we calculated the strength of the Swan band. The integral (swan_integ in Table A1) is derived between 4730 and 4738 Å. It should not be used as a substitute for a carbon abundance measurement, but only to sort the detections of carbon enhanced stars by their perceivable strength of the Swan band.

With so many spectra in our dataset, unexpected reduction and analysis problems can hinder the selection of carbon enhanced stars. In the first iteration, the results were ordered only by the value of the integrated Swan band region, but this proved to select too many spectra with reduction problems. Most of the problematic detections were caused by the incorrect normalization of spectra with strong, non-carbon molecular bands. This is best observable at the borders of a spectral range, where Swan bands are located in the case of HERMES spectra. There, normalization can be poorly defined in the case of numerous nearby absorption lines. In order to prevent miss-detections, additional limits on the shape ($S \ll 1$) and amplitude ($A \ll 1$) of the logΓ distribution were used to filter out faulty fitting solutions. Figure A6 represents one such example where the function was fitted to the absorption lines of a double-lined spectroscopic binary, producing a shape of the function that is not characteristic for the analysed molecular band head. We are also analysing the slope of the spectral difference and its integral in the limits of the Swan bands to remove spectra with reduction problems or peculiarity that would result in wrongly determined strength of the Swan band. One of the spectral trends we are trying to catch with those indicators is shown in Figure A6, where spectral difference and its linear fit are steeply rising at the border of spectrum.

By visual inspection of the algorithm diagnostic plots shown in Figure 1, we limited a final selection to 400 spectra with the strongest carbon enhancement that was still visually recognizable. The last selected spectrum is shown in Figure A4. With the selection of spectra with lower enhancement, we would risk to introduce possibly wrong classification of stars whose enhancement is driven by spectral noise levels.

3.2 Unsupervised classification

With numerous spectra of different stellar types, chemical composition and carbon enhancement, some of them might show different carbon features or be insufficiently distinctive to be picked out by the above supervised algorithm.

Another analysis technique, which is becoming increasingly popular is a dimensionality reduction procedure named t-distributed Stochastic Neighbor Embedding (t-SNE) (van der Maaten & Hinton 2008), that has already proved to be beneficial in comparison and sorting of unknown spectral features of the same dataset (Traven et al. 2017). This is done by projecting the complete spectra onto a 2D plane by computing similarities between all pairs of investigated spectra. It has been shown that the algorithm can cluster and distinguish spectra with absorption or emission features. The algorithm arranges spectra in a 2D plane, such that it clusters similar spectra together based on their similarity measure. As the transformation is variable and non linear, the actual distance between two objects in a final 2D plane does not linearly depend on the spectral similarity measure. This property of the t-SNE algorithm ensures
**Figure 2.** t-SNE projection of 588,681 observed spectra ranging between 4720 and 4890 Å. Red dots (756 spectra) mark a clump in the projection that was manually selected to contain carbon enhanced spectra. Superimposed blue dots represent carbon enhanced spectra determined by the supervised algorithm. Outside the t-SNE selected clump we have 224 spectra that were determined to be carbon enhanced only by the supervised method. All other analysed spectra are shown in gray shades, depending on their density in the 2D projection. Two circles indicate regions where the majority of CEMP candidates is located in the projection.

better coverage of the 2D plane in comparison to other dimensionality reduction methods.

The t-SNE projection shown in Figure 2 was computed from normalized spectra between 4720 and 4890 Å. To maximize the number of analysed spectra, no other limiting cuts than the validity of the wavelength solution (bit 1 in red_flag set to 0 by reduction pipeline, Kos et al. 2017) in this arm was used. This resulted in 588,681 individual spectra being analysed by the automatic unsupervised algorithm. This is 30k more spectra than in the case of supervised classification where we applied more strict criteria for the selection of spectra to be analysed (Section 3.1).

Without any prior knowledge about the location of objects of interest in the obtained projection, we would have to visually analyse every significant clump of stars in order to discover whether the carbon enhanced population is located in one of them. This can be simplified by adding the results of the supervised classification into this new projection. In Figure 2, the stars identified by the supervised classification are shown as blue dots plotted over gray dots representing all spectra that went into the analysis. The majority of blue dots are located in a clump on the left side of the projection. A high concentration of objects detected by a supervised method lead us to believe, that this isolated clump represents carbon enhanced objects in the t-SNE projection. We defined a polygon manually to select these stars.

Inspection of other blue labeled spectra outside the main clump revealed that their slight carbon enhancement could not be identified by the t-SNE similarity metric or the spectra comparison was dominated by another spectral feature.

Additional exploration of the t-SNE projection revealed two smaller groups of metal poor carbon enhanced spectra located inside circles shown in Figure 2. A confirmation that those regions are populated with metal poor stars can be found in Figure A1 where the spectra in projection are colour coded by [Fe/H] and \( T_{\text{eff}} \). To maximize the number of those objects in the published catalog, we manually checked all
undetected spectra in the vicinity of the detected ones. This produced an additional 13 detections.

3.2.1 t-SNE limitation

While checking the local neighborhood of some of the blue dots in Figure 2 that are strewn across the t-SNE projection we identified a possible limitation of our algorithm for automatic detection of specific peculiar spectra if their number is very small compared to the complete dataset. Figure 3 shows a collection of a few carbon enhanced spectra embedded between other normal spectra that were taken out of the right circled region in Figure 2. As they are quite different than the others they were pushed against the edge of a larger cluster in the projection, but their number is not sufficient to form a distinctive group of points in the projection. Therefore any automatic algorithm that would try to distinguish those objects based on a local density of points would most probably fail.

Another specific of the t-SNE projection that we must be aware of is how it computes the similarity between analyzed spectra. Combined similarity, which is computed as a sum of per pixel differences, has zero knowledge about the location where in the spectrum those differences occur. The red spectrum in Figure 4 with a slight signature of carbon enhancement in the range between 4734 Å and 4737 Å has been placed among spectra with similar physical properties. Its slight carbon enhancement and comparable spectral noise to other spectra in its vicinity are probably the reason why it was placed in such a region of the t-SNE projection. This could be solved by using a smaller portion of the spectrum in a dimensionality reduction. Although at the same time we might loose some other vital information about the star.

4 CANDIDATE CHARACTERISTICS

The final list of detected carbon enhanced stars consists of 918 stars, corresponding to 993 spectra detected by at least one of the described methods. Out of these, 63 stars were observed and identified at least twice and up to a maximum of four times. Those identifications belong to repeated observations that were performed at different epochs. Because not all of the observed spectra were considered in the classification procedure (because of the limitations described in Section 3) this is not the final number of stars with repeated observations. By searching among the complete observational dataset, the number of carbon enhanced stars with repeated observations increases to 90.

Out of those 90 stars, every repeated observation of 56 stars was classified as being carbon enhanced. In total we detected 76.5% of the carbon enhanced spectra among repeated observations where at least one of the repeats have been classified as having enhanced carbon features in its spectrum. The unclassified instances usually have a low SNR value that could decrease their similarity value towards other carbon enhanced stars in the t-SNE analysis or have incorrect stellar parameters and therefore be compared to an incorrect set of spectra during the supervised analysis.

4.1 Radial velocity variations

With repeated observations in the complete observational dataset, we can look into measured radial velocities and investigate a number of possible variables that should be high for certain types of carbon enhanced objects (Sperauskas et al. 2016). Taking into account all of the repeated observations in our dataset and not just the repeats among the identified spectra, 52 out of 90 stars show a minimum velocity change of 0.5 km s$^{-1}$ (70 stars with minimum change of 0.25 km s$^{-1}$) and a maximum of 45 km s$^{-1}$ in different time spans ranging from days to years. The detailed distribution is presented by Figure 5. That kind of change can hint at the presence of a secondary member or intrinsic stellar pulsation (Bergeat et al. 2002; Lloyd Evans 2010) because carbon enhanced stars are found among all long period variable classes (Mira, SRa, SRb). Follow-up observations are needed to determine their carbon subclass and subsequently the reason behind variations of radial velocity.

Figure 3. A collection of spectra that were determined to be mutually very similar by the t-SNE algorithm. Out of 46 spectra inside the right black circle in Figure 2 we identified 8 carbon enhanced spectra with visually very different and distinctive spectrum in the region from 4734 Å to 4737 Å that is also depicted in this plot. For easier visual recognizability, they are coloured in red.

Figure 4. Comparison between one of the detected carbon enhanced stars in red and its 30 closest neighbors in the t-SNE projection shown as black curves. Enhancement in the spectrum was probably not sufficiently distinct and dominated by the spectral noise as the spectrum was placed among other physically similar spectra without visible enhancement.
Visual inspection of variable candidates revealed that none of them shows obvious multiplications of absorption spectral lines, a characteristic of a double-lined binary system. Therefore we can conclude that none of them are binary members in which both components are of comparable luminosity and difference between their projected radial velocities is high enough to form a double-lined spectrum. From our simulations with median spectra, such line splitting becomes evident at the velocity difference of $\sim 14 \text{ km s}^{-1}$. If the components do not contribute the same amount of flux, the minimal difference increases to $\sim 20 \text{ km s}^{-1}$.

Any dwarf carbon enhanced star (dC) that exhibits enhancement of C\textsubscript{2} is thought to be or have been a member of a binary system in order to explain the presence of C\textsubscript{2} features in their spectra (Margon et al. 2018). Chemically enhanced material is thought to be accreted from the evolved AGB companion. Only two such systems have been identified spectroscopically to date (Dearborn et al. 1986; Margon et al. 2018), giving us the possibility to greatly increase the list with every additional confirmed object. The only detected dC star (for criteria see Section 4.2) with repeated observations shows that its radial velocity is unchanged on the order of $0.1 \text{ km s}^{-1}$ during the 2 years between consecutive observations. Hence, it can not be classified as a possible binary system from those two observations alone.

### 4.2 Stellar parameters

For the analysis of stellar parameters we used values as determined by The Cannon data interpolation method that was trained on actual observed HERMES spectra. To exclude any potentially erroneous parameter, we apply strict a flagging rule of flag\textsubscript{cannon}=0 (an extensive description of flagging procedure is found in Buder et al. 2018). Thus obtaining a set of 347 objects with trustworthy stellar parameters. Such a large percentage of flagged objects could be attributed to their nature as additional elemental enhancement that we are looking for might not be a part of the training set. A raised quality flag would hint that the spectrum is different from any other in the training set or that the fit is uncertain with a large $\chi^2$. Therefore flagged parameters have to be used with care, especially on the border of, and outside, the training set.

The majority (338) of the unflagged detected objects are giants and only 9 are confirmed to be dwarf stars based on their spectroscopic stellar parameters (Figure 6).

### 4.3 S-process elements

Focusing on a spectral signature of detected object inside and outside the t-SNE selected clump (Figure 2) we can further investigate which spectral feature might have contributed to their separation. The distributions of their abundances in Figure 7 and strength of spectral features corresponding to the same elements in Figure A7 hints to an enhancement of s-process elements among stars inside the selected clump. This additional enhancement might be another reason, besides the carbon enhancement, for the algorithm to cluster all of those stars as being different than the majority of spectra.

### 4.4 Lithium abundance

The derivation of elemental abundances for known carbon enhanced stars have showed that some of them can exhibit strongly enhanced levels of Li in their atmosphere (Abia et al. 1991). Lithium is thought to be produced by hot-bottom burning and brought to the surface from the stellar interior. Investigation of the Li line at 6707 Å revealed 32 such stars. Their spectra, centered around the Li feature, show greatly varying degree of absorption in Figure 8.

### 4.5 Subclasses

Following revision of the original MK classification (Keenan & Morgan 1941) introduced by Barnbaum et al. (1996), carbon stars are found in five different classes named C-H, C-R, C-J, C-N, and Barium stars. Of all the spectral indices proposed for the spectral classification we are only able to measure a small part of Swan C\textsubscript{2} bands and Ba II line at 6496 Å. For a more detailed classification of detected objects into...
Figure 7. Distribution of s-process element abundances for stars in three different groups. The most enhanced group in green represent carbon enhanced stars located in the t-SNE selected clump of stars. The red distribution presents all other detections that are placed around the projection and outside the clump. As a control group, the same distribution in black is shown for their closest t-SNE neighbors, therefore the black and red distribution contain equal number of objects. No abundance quality flags were used to limit abundance measurements.

Figure 8. Spectral subset of 32 lithium rich carbon enhanced stars among the identified stars. The highlighted wavelength region is used by The Cannon to determine the lithium abundance of a star.

proposed classes, we would need to carry out additional observations with a different spectroscopic setup to cover all the significant features.

Additionally, the features caused by the $^{13}\text{C}^{12}\text{C}$ molecule are strongly enhanced only for a handful of spectra in our dataset, therefore we did not perform any isotopic ratio analysis or identification of possible C-J objects, which are characterized by strong isotopic Swan bands.

According to the abundance trends presented in Section 4.3 and the classification criteria defined by Barnbaum et al. (1996) we could argue that the stars selected from the t-SNE projection belong to the C-N subclass. Their s-process elements are clearly enhanced over Solar values (Figure 7), but the actual values should be treated with care as they are mostly flagged by The Cannon. This uncertainty might come from the fact that the training set does not cover carbon enhanced stars and/or stars with such enhancement of s-process elements.

4.6 Match with other catalogs

In the literature we can find numerous published catalogs of carbon enhanced (CH) stars (Christlieb et al. 2001; Alksnis et al. 2001; Ji et al. 2016) and CEMP stars (Komiya et al. 2007; Masseron et al. 2010; Placco et al. 2010, 2014; Abate et al. 2015; Yoon et al. 2017) observed by different telescopes and analysed in an inhomogeneous way. Most of those analyses were also performed on a spectra of lower resolving power than HERMES, therefore some visual differences are expected for wide molecular bands. By matching published catalogs with GALAH observations that were analyzed by our procedures, we identified 44 stars that matched with at least one catalog. Of these, 28 were found in CH catalogs and 16 CEMP catalogs.

From the stars recognized as CEMPs in the literature we were able to recover only 1 of them. Visual assessment of the diagnostic plots provided by our analysis pipeline proved that the remaining 15 CEMP matches expresses any observable carbon enhancement in Swan bands and were therefore impossible to detect with the combination of our algorithms. The reason for this difference between our and literature results might be in the CEMP selection procedure employed by the aforementioned literature. Every considered study selects their set of interesting stars from one or multiple literature sources based on values of $\left[\text{M}/\text{H}\right]$ and $\left[\text{C}/\text{Fe}\right]$ that were measured from the atomic spectral lines and not molecular lines.

The match is larger in the case of CH matches where we were able to confirm 11 out of 33 possible matched carbon enhanced stars. As the observed molecular bands are prominent features in the spectra, we explored possible reasons for our low detection rate. Visual inspection of spectra for the remaining undetected matched stars proved that they also show no or barely noticeable carbon enhancement in Swan bands, therefore reason must lie in the detection procedure used in the cited literature. Christlieb et al. (2001) used low resolution spectra to evaluate enhancement of C$_2$ and CN bands. The results are also given in their electronic table. In here, all of our undetected stars are marked to contain CN enhancement but no C$_2$. Combining this with Figure 9 we speculate that those stars occupy a narrow range of parameter space where C$_2$ is not enhanced and therefore undetectable in HERMES spectra.

The position of all stars matched with the literature is
5 METAL POOR CANDIDATES

CEMP stars are defined in the literature as having low metallicity \([\text{M/H}] < -1\) and strong carbon enrichment \([\text{C/Fe}] > +1\). In the scope of this analysis we assume that our measurement of \([\text{Fe/H}]\) is a good approximation for the metallicity. To be sure about this we compared \([\text{M/H}]\) values of CEMP stars found in the literature and \([\text{Fe/H}]\) derived by The Cannon for the same stars. The relation between them is shown in Figure 10. We see that our values start deviating from the published values at metallicities below \(-1.5\) dex. Below that threshold the difference are in the range of \(-1\) dex, but the same trend is obvious for both datasets. The uncertainty of the published \([\text{M/H}]\), derived from multiple sources, can reach up to 0.5 dex.

Taking unflagged The Cannon parameters and abundances of the detected objects we can determine possible CEMP candidates among our sample. As also shown by Figure 11 our set of carbon-enhanced stars consists of 41 objects with \([\text{Fe/H}] < -0.5\) and 2 objects with \([\text{Fe/H}] < -1.0\). If we also include potentially incorrect parameters, the number of objects with \([\text{Fe/H}] < -1.0\) increases to 28, which is equal to 2.8% of detected carbon enhanced spectra. In any case, none of them have a valid carbon abundance determination. Analyzing HERMES spectra to determine carbon abundance is difficult because the automatic analysis is based on only one very weak atomic absorption line that is believed to be free of any blended lines. Consequently, we are also not able to measure the \([\text{C/O}]\) abundance ratio as majority of determined \([\text{C/Fe}]\) abundances is flagged as unreliable. Another methodology is needed to determine the abundance and confirm possible CEMP candidates.

When looking at the distribution of \([\text{Fe/H}]\) for the complete set of observed stars, we find a comparable distribution as for carbon enhanced stars. Similarly, about 1.8% of stars are found to be metal poor with \([\text{Fe/H}] < -1.0\).
6 FOLLOW-UP OBSERVATION

To further classify and analyse one of the detected objects, a star with 2MASS identifier J11333341-0043060 was selected for follow-up observation. We acquired a high resolution Echelle spectrum (with the resolving power R \( \sim 20,000 \)) using a spectrograph mounted on the 1.82 m Copernico telescope located at Cima Ekar (Asiago, Italy). Because only few of our detected candidates are observable from the Asiago observatory, the selected target was chosen to be a metal poor CEMP candidate with an iron abundance [Fe/H] of \(-0.96\) as determined by The Cannon. The selected star, with \( V = 12.79 \), was on the dark limit of the used telescope, therefore low SNR was expected. The one hour long exposure of the selected object was fully reduced, normalized order by order and shifted to the rest frame.

Although the acquired spectrum covers a much wider and continuous spectral range (from 3900 to 7200 Å) than HERMES spectra, only subsets, relevant for the classification of carbon enhanced stars are presented in Figure 12. They were identified by visually matching our observed spectrum with the published moderate resolution spectral atlas (Barnbaum et al. 1996) of peculiar carbon stars. Where available, the GALAH spectrum is shown alongside the Asiago spectrum. Carbon enhancement is not expected to vary over a period of several years, so both spectra should show similar features. The second and fourth panel in Figure 12 confirm that both observations indicate a similar degree of carbon enhancement.

Following the classification criteria of carbon stars, we determined that the star belongs to the C-H subclass. The definitive features for this class are strong molecular CH bands, prominent secondary P-branch head near 4342 Å (top panel in Figure 12) and noticeable Ba II lines at 4554 and 6496 Å (Li et al. 2018), which are all observed. The star definitely does not have a high ratio between \(^{13}\)C and \(^{12}\)C isotopes as the Swan features corresponding to \(^{13}\)C are not clearly present, therefore it can not be of a C-J subclass.

Following the current state of knowledge (McClure & Woodsworth 1990; Jorissen et al. 2016; Sperauskas et al. 2016) that most, if not all, C-H stars show clear evidence for binarity, we compared the radial velocity between both observations. They hint at the variability of the object as the follow-up radial velocity \((126.75 \pm 1.63 \text{ km s}^{-1})\) deviates by more than \(3 \text{ km s}^{-1}\) from the velocity \((123.43 \pm 0.08 \text{ km s}^{-1})\) observed as part of the GALAH survey. The time span between the two observations is more than 2.5 years, where the exact JD of the observation is 2458090.702 for the Asiago spectrum and 2457122.095 for the GALAH spectrum. Further observations along the variability period would be needed to confirm whether it is a multiple stellar system.

7 CONCLUSIONS

This work explores stellar spectra acquired by the HERMES spectrograph in order to discover peculiar carbon enhanced stars, which were observed in the scope of multiple observing programs conducted with the same spectrograph.

We show that the spectra of such stars are sufficiently different from other stellar types to be recognizable in high resolution spectra with limited wavelength ranges. This can
be done using a supervised procedure where some knowledge about the effects of carbon enhancement of observed spectra is put into the algorithm or using an unsupervised method. The latter was used to identify observed stars solely on the basis of acquired spectra. By combining both methodologies we identified 918 unique stars with evident signs of carbon enhancement of which 12 were already reported in the literature. Out of all matched objects from the literature we were unable to detect and confirm 16 (57%) CH and 15 (93%) CEMP stars with our procedures. As some of those objects were proven to contain carbon enhancement detectable outside the HERMES wavelength ranges, this would have to be taken into account to say something about the underlying population of carbon enhanced stars.

The identified objects were then separated into dwarf and giant populations using their stellar atmospheric parameters that were also used to select possible CEMP candidates. All of the detections with multiple observations at different epochs were investigated for signs of variability.

With a follow-up observation of one of the identified stars, we were able to confirm existence of carbon rich molecules in its atmosphere in a wider wavelength range. This was used to determine its subclass. Variation in radial velocity points to a possible variable nature of the star or binarity that is common for C-II stars.

Follow-up observations are required to confirm variability of radial velocities observed for some of the detected carbon enhanced stars and further investigate their nature. Careful spectral analysis, with inclusion of carbon enhancement in models, is needed to confirm the metallicity levels of the metal poor candidates.

In the electronic table of the paper, we have included a list of the detected stars. The list also includes stars from the literature, matched with our observations, for which we were unable to confirm their carbon enhancement. The list could be used to plan further observations, allowing a better understanding of these objects.

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## Table A1

List and description of the fields in the published catalog of detected objects and objects matched with multiple literature sources.

| Field          | Unit | Description                          |
|----------------|------|--------------------------------------|
| source_id      |      | Gaia DR2 source identifier           |
| sobject_id     |      | Unique internal per-observation star ID |
| ra             | deg  | Right ascension from 2MASS, J2000     |
| dec            | deg  | Declination from 2MASS, J2000         |
| det_sup        | bool | Detected by supervised fitting method |
| det_usup       | bool | Detected by t-SNE method              |
| swan_integ     |      | Swan band strength if determined      |
| teff           | K    | The Cannon effective temperature $T_{\text{eff}}$ |
| e_teff         | K    | Uncertainty of determined $T_{\text{eff}}$ |
| logg           | dex  | The Cannon surface gravity log $g$    |
| e_logg         | dex  | Uncertainty of determined log $g$     |
| feh            | dex  | The Cannon iron abundance $\left[\text{Fe}/\text{H}\right]$ |
| e_feh          | dex  | Uncertainty of determined $\left[\text{Fe}/\text{H}\right]$ |
| flag_cannon    | int  | The Cannon flags in a bitmask format |
| type           |      | G for giants and D for dwarfs         |
| rv_var         | bool | Is radial velocity variable           |
| li_strong      | bool | Shows strong lithium absorption       |
| cemp_cand      | bool | Is star CEMP candidate                |
| bib_code       |      | ADS bibcode of the literature match   |

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## APPENDIX A: TABLE DESCRIPTION

In the Table A1 we provide a list of metadata available for every object detected using the methodology described in this paper. The complete table of detected objects and its metadata is available only in electronic form at the CDS.
**Figure A1.** Spatial distribution of all available measurements of $T_{\text{eff}}$ (left panel) and $[\text{Fe}/\text{H}]$ (right panel) as determined by *The Cannon*. Dots, representing analysed spectra in the t-SNE projection, are colour coded by their parameter values. Colours and their corresponding values are explained by a colourbar under the graph.

**Figure A2.** t-SNE projection with marked known carbon enhanced and CEMP objects from multiple different catalogs found in the literature that are also part of our analysed set of spectra.
Figure A3. Equivalent plot as in the Figure 1 but presenting an example of a metal rich star with multiple strong Swan features around 4737 Å and 4745 Å. Presented star has a 2MASS identifier J13121354-3533120 and it is known Galactic carbon star.

Figure A4. Equivalent plot as in the Figure 1 showing the last of 400 spectra, ordered by their degree of carbon enhancement, selected by the supervised methodology.
**Figure A5.** Equivalent plot as in the Figure 1 but representing grossly over exaggerated carbon enhancement by a fit stuck in a reduction problem (a cosmic ray in a subtracted sky spectrum).

**Figure A6.** Equivalent plot as in the Figure 1 but representing a fit to absorption lines of resolvable spectroscopic binary. Final fit is not skewed as would be expected in the case of carbon enhancement.
Figure A7. Spectral subset around the absorption features in the blue arm that were used to determine abundances of Fe and s-process element abundances. Same colour coding is used as in Figure 7. Spectra inside the t-SNE determined clump are shown in red, and outside it in green. Median of all spectra is shown with a bold line of the same colour. The shaded area gives the actual range considered for the computation of abundances.