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GRB host galaxies with strong H$_2$ absorption: CO-dark molecular gas at the peak of cosmic star formation

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ABSTRACT

We present a pilot search of CO emission in three H$_2$-absorbing, long-duration gamma-ray burst (GRB) host galaxies at $z \sim 2$–3. We used the Atacama Large Millimeter/submillimeter Array (ALMA) to target the CO(3–2) emission line and report non-detections for all three hosts. These are used to place limits on the host molecular gas masses, assuming a mass-metallicity-dependent CO-to-H$_2$ conversion factor ($\alpha_{\text{CO}}$). We find, $M_{\text{mol}} < 3.5 \times 10^{10}$ M$_\odot$ (GRB 080607), $M_{\text{mol}} < 4.7 \times 10^{11}$ M$_\odot$ (GRB 120815A), and $M_{\text{mol}} < 8.9 \times 10^{11}$ M$_\odot$ (GRB 181020A). The high limits on the molecular gas mass for the latter two cases are a consequence of their low stellar masses $M_*$ ($M_* \lesssim 10^8$ M$_\odot$) and low gas-phase metallicities ($Z \sim 0.03$ Z$_\odot$). The limit on the $M_{\text{mol}}/M_*$ ratio derived for GRB 080607, however, is consistent with the average population of star-forming galaxies at similar redshifts and stellar masses. We discuss the broader implications for a mass-metallicity-dependent CO-to-H$_2$ conversion factor and demonstrate that the canonical Galactic $\alpha_{\text{CO}}$ will severely underestimate the actual molecular gas mass for all galaxies at $z > 1$ with $M_* < 10^9$ M$_\odot$.

To better quantify this we develop a simple approach to estimate the relevant $\alpha_{\text{CO}}$ factor based only on the redshift and stellar mass of individual galaxies. The elevated conversion factors will make these galaxies appear CO-‘dark’ and difficult to detect in emission, as is the case for the majority of GRB hosts. GRB spectroscopy thus offers a complementary approach to identify low-metallicity, star-forming galaxies with abundant molecular gas reservoirs at high redshifts that are otherwise missed by current ALMA surveys.

Key words: gamma-ray burst: general – ISM: molecules – galaxies: high-redshift – galaxies: ISM – galaxies: star formation.

1 INTRODUCTION

Since long-duration gamma-ray bursts (GRBs) are linked to the death of massive stars (Hjorth et al. 2003; Stanek et al. 2003; Woosley & Bloom 2006; Cano et al. 2017), they are expected to trace star formation through cosmic time (Wijers et al. 1998; Kistler et al. 2009; Robertson & Ellis 2012; Greiner et al. 2015). GRB-selected galaxies therefore probe the underlying population of star-forming galaxies that are not biased towards the most luminous and massive galaxies unlike traditional emission-selected galaxy surveys. Moreover, the short-lived optical afterglows following GRBs are so bright that the plethora of absorption features that are imprinted from the interstellar medium (ISM) of the GRB host on the afterglow spectrum can be studied in detail (e.g. Jakobsson et al. 2004; Prochaska et al. 2007; Vreeswijk et al. 2007; Fynbo et al. 2009). After the first few afterglow spectra were obtained it was clear that GRB-host absorption systems typically probe sightlines with the highest H$_1$ column densities of the so-called damped Ly$\alpha$ absorbers (DLAs; Vreeswijk et al. 2004; Jakobsson et al. 2006; Fynbo et al. 2009), related to their small impact parameters. DLAs provide the most effective and detailed probe of neutral gas in high-redshift galaxies and contain most of the neutral gas at high redshift (Noterdaeme et al. 2009). Given their direct link to star formation and the very high column densities of gas typically detected in GRB afterglow spectra, the low detection rate of molecular hydrogen H$_2$ (from the UV Lyman–Werner bands) was initially a puzzle (e.g. Tumlinson et al. 2007; Ledoux et al. 2009). The first detection of H$_2$ in a GRB absorber was observed in the remarkable afterglow spectrum of GRB 080607 (Prochaska et al. 2009). Since then, eight more H$_2$-bearing GRB absorbers have been detected (Krühler et al. 2009).

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$^{1}$Though see also the tentative detection of H$_2$ in the afterglow of GRB 060206 reported by Fynbo et al. (2006).
2 SAMPLE AND OBSERVATIONS

2.1 GRB hosts with strong H$_2$ absorption

We targeted the three GRB-hosts with strongest H$_2$ absorption known to date, all at z > 2. These are: GRB 080607 (Prochaska et al. 2009), GRB 120815A (Krührer et al. 2013), and GRB 181020A (Heintz et al. 2019). All show H$_2$ column densities above $N$(H$_2$) = 10$^{20}$ cm$^{-2}$ and also constitute the GRB absorbers with the largest molecular-hydrogen fractions, $f_{H_2}$ (Bolmer et al. 2019; Heintz et al. 2019). GRB 080607 also shows a high absorption-derived gas-phase metallicity consistent with Solar ([X/H] = -0.2) and significant dust extinction $A_V$ ~ 3 mag. On the contrary, GRBs 120815A and 181020A show relatively low metallicities of $[X/H]$ ~ -0.3 and modest extinction in the line of sight ($A_V$ = 0.2 – 0.3 mag).

The host-galaxy emission counterpart of GRB 080607 is well detected in several bands (Chen et al. 2010; Wang, Chen & Huang 2012). Here, we adopt the stellar population properties derived by Corre et al. (2018), inferring log ($M_*/M_\odot$) = 10.45 ± 0.10. For the other two, more recently detected GRBs, no host-galaxy counterpart has been identified yet. For these, we instead rely on previous work connecting DLAs to their emission counterparts (Møller et al. 2013; Neelam et al. 2013; Christensen et al. 2014). Following Arbasalmani et al. (2015), we assign an impact parameter of 2.3 kpc and compute the predicted stellar masses based on the prescription by Christensen et al. (2014, their equation 3). This yield stellar masses of log ($M_*/M_\odot$) = 7.9 ± 0.4 (GRB 120815A) and log ($M_*/M_\odot$) = 7.8 ± 0.4 (GRB 181020A), where the uncertainties are dominated by the internal scatter in log $M_*$.

3 RESULTS

We searched both the 25 and 100 km s$^{-1}$ channel width ALMA spectral data cubes for emission originating from CO(3 – 2) at the relevant redshifts for the GRBs in our sample. No line emission is detected at the position of any of the GRBs (but see, e.g. Neelam et al., in preparation, for a detection of CO(2 – 1) from the strong intervening Mg II absorber towards GRB 120815A). We also did not detect continuum emission at the positions of any of the three GRBs, but there appears to be continuum emission from an unrelated galaxy in the field of GRB 120815A.

We extracted 1D spectra from the ALMA spectral data cube centred on the positions of the GRB afterglows or host galaxies, shown in Fig. 1. We derived 3σ upper limits on the velocity-integrated flux densities of $<0.09$ Jy km s$^{-1}$ (GRB 080607), $<0.018$ Jy km s$^{-1}$ (GRB 120815A), and $<0.016$ Jy km s$^{-1}$ (GRB 181020A), assuming line widths for the emission-line profiles of FWHM $= 300$ km s$^{-1}$ (GRB 080607) and FWHM $= 50$ km s$^{-1}$ (GRBs 120815A and 181020A), appropriate for galaxies in their given mass ranges (e.g. Tiley et al. 2016). We then derived the corresponding CO(3 – 2) line luminosities (following equation 3 from Solomon & Vanden Bout 2005) of $L_{CO(3–2)} < 4.56 \times 10^9$ K km s$^{-1}$ pc$^2$ (GRB 080607), $L_{CO(3–2)} < $
GRBs 080607, 120815A, and 181020A, respectively. All the values for star-forming galaxies; Dessauges-Zavadsky et al. 2015) and adopting a metallicity-dependent CO-to-H2 conversion factor:

\[ r \]

into total molecular gas masses assuming a line ratio of \( \alpha \) following Heintz & Watson 2020. This \( \alpha_{\text{CO}} \)-metallicity relation is the average between the locally derived (Israel 1997; Leroy et al. 2013; Amorín et al. 2016) and high-redshift (Genzel et al. 2012) inferred relations, and is calibrated to the molecular gas mass evolution of \( M_\star > 10^{10} M_\odot \) main-sequence star-forming galaxies, \( M_{\text{mol}}/M_\star = 0.1 \times (1+z)^2 \) (Geach et al. 2011; Carilli & Walter 2013).

While the ratio of molecular gas to stellar mass of the hosts of GRBs 120815A and 181020A is poorly constrained, the host galaxy of GRB 080607 is only marginally consistent with that expected for typical star-forming galaxies at similar redshifts and stellar masses. This is in stark contrast with the population of DLAs at \( z \geq 2 \) observed in quasar sightlines, shown as the green symbols in Fig. 2, which overall show a significant excess of molecular gas (Neeleman et al. 2018; Kanekar et al. 2020). Assuming a lower CO excitation of \( r_{11} = 0.4 \) (mostly representative of the \( z \approx 2 \) population; Boogaard et al. 2020), however, results in a consistent limit on the molecular-to-stellar mass content. The high upper limits on the \( M_{\text{mol}}/M_\star \) fractions of the hosts of GRBs 120815A and 181020A are mainly due to their low metallicities requiring high CO-to-H2 conversion factors to infer their molecular gas content. These host galaxies, however, also have lower stellar masses than the typical star-forming galaxies probed in CO at similar redshifts (e.g. Tacconi et al. 2013, 2018). In the next section, we will explore these particular high-redshift, low-mass, low-metallicity hosts.

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**Figure 1.** CO flux density as a function of velocity, where \( v_{\text{rel}} = 0 \, \text{km} \, \text{s}^{-1} \) corresponds to the redshift of the absorbers, \( z_{\text{GRB}} \). None of the spectra show detection of CO(3 – 2) lines at the redshifts of the GRBs.

5.51 \times 10^{10} \, \text{K} \, \text{km} \, \text{s}^{-1} \, \text{pc}^2 \) (GRB 120815A), and \( L_{\text{CO(3–2)}}^{\text{rel}} < 7.12 \times 10^{9} \, \text{K} \, \text{km} \, \text{s}^{-1} \, \text{pc}^2 \) (GRB 181020A).

We converted the measured CO(3 – 2) line luminosities into total molecular gas masses assuming a line ratio of \( r_{31} = L_{\text{CO(3–2)}}/L_{\text{CO(1–0)}} = 0.57 \) (which is the observed average for \( z > 1 \) star-forming galaxies; Dessauges-Zavadsky et al. 2015) and adopting a metallicity-dependent CO-to-H2 conversion factor:

\[ \alpha_{\text{CO}}(Z) = 4.5 \times (Z/Z_\odot)^{-1.40} M_\odot \, (\text{K} \, \text{km} \, \text{s}^{-1} \, \text{pc}^2)^{-1}, \]

(1)

following Heintz & Watson 2020. This \( \alpha_{\text{CO}} \)-metallicity relation is the average between the locally derived (Israel 1997; Leroy et al. 2011; Bolatto, Wolfire & Leroy 2013; Amorín et al. 2016) and high-redshift (Genzel et al. 2012) inferred relations, and is calibrated to galaxies at \( z > 1 \). This yields upper limits on the molecular gas masses of \( \log (M_{\text{mol}}/M_\star) < 10.54, 11.67, \) and 11.95, for the hosts of GRBs 080607, 120815A, and 181020A, respectively. All the values derived in this section are summarized in Table 1.

In Fig. 2, we show the upper limits of the molecular gas-to-stellar mass ratio of the three GRB hosts. In each case, the molecular gas mass ratio is determined by using either the above \( \alpha_{\text{CO}} \)-metallicity relation and the host absorption-derived metallicities, or by assuming a constant MW-like conversion factor of \( \alpha_{\text{CO,MW}} = 4.3 \, M_\odot \, (\text{K} \, \text{km} \, \text{s}^{-1} \, \text{pc}^2)^{-1} \) (Bolatto et al. 2013). For comparison, we overplot the recent compilation of CO observations of GRB host galaxies from Hatsukade et al. (2020), spanning the redshift range \( z = 0.0–2.5 \). In addition, we show the track that broadly characterizes the molecular gas mass evolution of \( M_\star > 10^{10} M_\odot \) main-sequence star-forming galaxies, \( M_{\text{mol}}/M_\star = 0.1 \times (1+z)^2 \) (Geach et al. 2011; Carilli & Walter 2013).

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**Table 1.** Sample properties of the H2-bearing GRB host galaxies.

| GRB   | \( z_{\text{GRB}} \) | \( \log N(\text{H}_2) \) (cm\(^{-2}\)) | \( \log f_{\text{HI}} \) | [X/H] | \( A_v \) (mag) | \( \log M_\star \) (\( M_\odot \)) | \( L_{\text{CO(3–2)}}^{\text{rel}} \) (10\(^6\) K km s\(^{-1}\) pc\(^2\)) | \( \log M_{\text{mol}} \) (\( M_\odot \)) |
|-------|----------------------|-----------------------------------|----------------|------|----------------|-------------------------------|----------------------------------|----------------------------------|
| 080607| 3.0363               | 21.20 ± 0.20                      | –1.23 ± 0.24  | –0.2 | 2.58 ± 0.45   | 10.45 ± 0.10                  | <4.56                            | <10.54                           |
| 120815A| 2.3582             | 20.42 ± 0.08                      | –1.39 ± 0.09  | –1.45 ± 0.03 | 0.19 ± 0.04   | 7.90 ± 0.40                  | <0.55                            | <11.67                           |
| 181020A| 2.9379              | 20.40 ± 0.04                      | –1.51 ± 0.06  | –1.57 ± 0.06 | 0.27 ± 0.02   | 7.80 ± 0.40                  | <0.71                            | <11.95                           |
and low-metallicity hosts in context to the underlying field-galaxy population.

4 DISCUSSION

4.1 Field environment of GRB hosts

Studies of absorption-selected galaxies in the line of sight towards bright background quasars have revealed that many of these DLAs are found in environments with other nearby galaxies, both at low (Kacprzak, Murphy & Churchill 2010; Rahmani et al. 2018) and high redshifts (Francis & Hewett 1993; Möller & Warren 1993; Fynbo et al. 2003). These DLA counterparts have mostly been detected based on strong rest-frame optical emission lines, or from Ly α emission from the galaxy counterparts, but are now also being increasingly detected in CO (Fynbo et al. 2018; Klitsch et al. 2018). With the data presented here we can examine in an independent way whether GRB-DLAs also appear to be part of larger galaxy complexes.

Each ALMA cube covers ≈45 arcsec subtended on the sky, which corresponds to 350 kpc and 375 kpc at the target redshifts (z = 3 and z = 2.3), respectively. We do not detect emission from CO(3 − 2) in galaxies within the ≈±1500 km s⁻¹ covered by the data cube, corresponding to z = ±0.02. We thus do not find evidence for galaxy clustering in this (albeit small) sample of CO-surveyed GRB hosts. The quasar-DLA bias towards CO-emitting galaxy groups could partly be explained by the preferential high-metallicity these systems were selected on. Consequently, absorption-line features in quasar sightlines could therefore be influenced by these more populous environments (e.g., Arabsalmani et al. 2020), rather than than tracing the line of sight through a single galactic disc (Fynbo et al. 2018). GRB sightlines could therefore provide a cleaner probe of absorption-derived quantities and correlations (e.g. Arabsalmani et al. 2015, 2018a).

4.2 Mass and redshift dependence of α_{CO}

One of the most promising aspects of studying the CO emission associated to GRB hosts, is that it allows us to probe the molecular gas content in low-metallicity galaxies which are otherwise missed by field-selected surveys. As metallicity decreases both with increasing redshift and decreasing galaxy mass, we expect α_{CO} to show a strong mass and redshift dependence. In fact, even a metallicity of Z/Z⊙ ≈ 10 per cent would imply an α_{CO} value more than an order of magnitude higher than the average Milky Way conversion factor α_{CO,MW}. The reason why this is still largely applied to high-redshift galaxies, is due to the difficulty in measuring the gas-phase metallicity in low-metallicity, low-luminosity galaxies. Below, we aim to improve on this and present a simple relation that conveniently expresses the CO-to-H₂ conversion factor, in addition to the M_{mol}/M_⋆ ratio, as a function of stellar mass and redshift.

Starting from the metallicity-dependent α_{CO} relation (equation 1), we can connect α_{CO} directly to the galaxy stellar mass at any given redshift via mass–metallicity relations. For the latter, Savaglio et al. (2005) provided a convenient fit, valid for a wide mass and redshift ranges, that has been improved upon by Genzel et al. (2015) who combined metallicity prescriptions from four different studies. The latter paper, however, introduces a solar-metallicity cut-off at low redshift, which is too limiting for our purposes. Instead, we start with the results of Maiolino et al. (2008), one of the four prescriptions used by Genzel et al. (2015), that provide at each redshift bin a mass–metallicity relation:

$$12 + \log(O/H) = -0.864(\log(M_\star) - \log(M_\odot))^2 + K_0.$$  (2)

Here, M_⋆ is the stellar mass and log M_⊙ and K_0 are constants in each bin, their values given in Maiolino et al. (2008, table 5). We fit a function of (1 + z) to these factors and find that log M_⋆ = 2.59 × log (1 + z) + 11.05 and K_0 = 8.9, provides a good representation of their redshift evolution. This mass-metallicity relation is shown in Fig. 3. The K_0-values can also be well fitted by a second-order polynomial in z, but this significantly underpredicts the metallicity at z > 3 compared to other calibrators. With the above form for log M_⋆ and a constant K_0, we allow for super-solar metallicities and find metallicity values that are bracketed by the mass–metallicity calibrations of Savaglio et al. (2005) and Genzel et al. (2015), at all but the lowest masses and lowest redshifts. We include an uncertainty of 0.2 dex in this mass–metallicity relation, representing the typical dispersion around this relation.

Connecting the α_{CO}–metallicity relation with the mass–metallicity calibration derived, we obtain a simple expression for α_{CO}:

$$\log(\alpha_{CO}) = 0.121(\log(M_\star) - \log(M_\odot))^2 + 0.359,$$  (3)

depending only on the stellar mass, M_⋆, of a galaxy at redshift z. Again, log M_⋆ = 2.59 × log (1 + z) + 11.05, representing the characteristic turn-over stellar mass at a given redshift. These relations allow us to estimate directly the appropriate molecular gas mass conversion factors for main-sequence star-forming galaxies at any given redshift and stellar mass for which the mass–metallicity calibration is accurate.

Based on this combined relation, we explore how α_{CO} evolves in typical star-forming galaxies as a function of redshift and stellar mass. We show these evolutionary tracks in Fig. 4 for a set of representative redshifts and stellar masses, namely z = 0, 1, 3, 5 and M_⋆ = 10^8, 10^9, 10^{10}, 10^{11} M⊙. The shaded error regions on each curve represent the combined errors from the typical dispersion around the mass–metallicity relation (σ_M2 = 0.2 dex) and the scatter in the best-fitting α_{CO}–metallicity relation. In the figure, we also
compare the evolution of $\alpha_{CO}$ to the average values observed in the Milky Way and in local galaxies (i.e. equivalent to solar metallicities) of $\alpha_{CO,MW} = 3.5 - 5.5 M_\odot (K\text{ km s}^{-1}\text{ pc}^2)^{-1}$ (Bolatto et al. 2013).

As suggested by the mass–metallicity relation (equation 2), the evolution of $\alpha_{CO}$ depends strongly on both galaxy mass and redshift. We find that the Galactic conversion factor $\alpha_{CO,MW}$ is a suitable approximation for massive galaxies, with $M_\star \sim 10^{11} M_\odot$, at all redshifts in the range $z = 0 - 5$. Similarly, galaxies at $z \sim 0$ with stellar masses in the range $M_\star = 10^9 - 10^{11} M_\odot$ also show conversion factors consistent with that observed in the Milky Way within the uncertainties. However, we do recover an overall increase in $\alpha_{CO}$ at decreasing stellar masses for fixed redshifts as expected. Galaxies with $M_\star < 10^9 M_\odot$ and also most galaxies at $z > 1$ (except the most massive ones, $M_\star \sim 10^{11} M_\odot$), i.e. the majority of GRB hosts, show $\alpha_{CO}$ values significantly higher than $\alpha_{CO,MW}$. For instance, a galaxy with $M_\star = 10^9 M_\odot$ at $z \sim 0$ will have a conversion factor of $\alpha_{CO} = 30.5 M_\odot (K\text{ km s}^{-1}\text{ pc}^2)^{-1}$, almost an order of magnitude higher than the average value observed in the Milky Way. At higher redshifts, where CO is now being increasingly detected (e.g. Tacconi et al. 2010, 2013, 2018; Walter et al. 2011; Decarli et al. 2016; Bothwell et al. 2017; Pavesi et al. 2018; Valentino et al. 2018, 2020; Aravena et al. 2019; Riechers et al. 2019; Harrington et al. 2021), even massive galaxies with $M_\star \sim 10^{10} M_\odot$ at $z = 2.5$ will have conversion factors of $\alpha_{CO} = 12.3 M_\odot (K\text{ km s}^{-1}\text{ pc}^2)^{-1}$, exceeding the average Galactic value by a factor of 2–3.

The redshift- and mass-dependent evolution of $\alpha_{CO}$ can now be used to shed light on the observed $M_{\text{mol}}/M_\star$-redshift relation. As discussed in the previous section, it has been shown that this ratio approximately follows $M_{\text{mol}}/M_\star \propto (1 + z)^{2.5}$, for star-forming galaxies (Tacconi et al. 2018). In the following analysis we assume a constant value of $L_{CO,1-0} = 2 \times 10^{10} \text{ K km s}^{-1}\text{ pc}^2$, which is the average luminosity of galaxies at $z \sim 1 - 3$ derived from the PHIBSS sample (corrected by $r_{31}$; Tacconi et al. 2013). We caution that the CO luminosity has been found to be strongly correlated with the stellar mass of each galaxy (e.g. Inami et al. 2020), and that this analysis should only be treated as a simple model to explore the evolution of $M_{\text{mol}}/M_\star$ with redshift.

In Fig. 5, we plot the resulting $M_{\text{mol}}/M_\star$-curves for a range of stellar masses, $M_\star = 10^6, 10^9, 10^{10}, 10^{11} M_\odot$, covering redshifts from $z = 0 - 3$, with the GRB hosts overplotted for reference. For the hosts of GRBs 120815A and 181020A, the upper limits are weaker than expected from their host stellar masses, whereas the limit on the molecular gas mass in GRB 080607 is close to that expected from its host stellar mass. Overall, most of the CO detections in GRB hosts do fall between the curves with $M_\star = 10^{10} - 10^{11} M_\odot$ (Hatsukade et al. 2020), as expected from their stellar masses, though some hosts appear to show significant molecular deficits (in particular at $z \lesssim 1$).

Finally, in Fig. 5, we show the expected evolution of $M_{\text{mol}}/M_\star$ evolution track assuming a 0.2 dex metallicity decrease per unit redshift (as observed for absorption-selected galaxies, e.g. Rafelski et al. 2012; De Cia

**Figure 4.** The evolution of the molecular gas mass to CO(1 – 0) luminosity ratio, $\alpha_{CO}$, for regular star-forming galaxies typical galaxy scaling relations. The left-hand panel shows the evolution as a function of redshift for selected galaxy stellar masses and the right-hand panel shows the corresponding evolution as a function of stellar mass for the given redshifts. The shaded error region on each curve represents the combined errors from the mass–metallicity relation and $\alpha_{CO}$–metallicity relation (see the main text for more details). The average value observed in the Milky Way is shown with the blue horizontal band in both panels.

**Figure 5.** Same as Fig. 4, but with curves of constant stellar mass overplotted. Only the upper limits using the $\alpha_{CO}$–metallicity relation (equation 1) are shown for the three GRB hosts from this work. All plotted data points are colour-coded according to stellar mass. In addition, the dotted curve shows the expected evolution for a 0.2 dex metallicity decrease per unit redshift, starting at $\approx 1.5 \times Z_\odot$. 

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et al. 2018). The trend start at $Z = 1.5Z_\odot$ at $z = 0$, and we again assume $L_{CO}(1-0) = 2 \times 10^{10}$ K km s$^{-1}$ pc$^2$ and the mass–metallicity and $M_{\text{CO}}$–mass relations described above (shown as the dotted line in the figure). It is clear that this evolution track approximately follows the relation, $M_{\text{mol}}/M_\star \propto (1+z)^{3.5}$, that has been considered to broadly describe the evolution of the more massive galaxies, $M_\star > 10^{10}$ M$_\odot$.

While the stellar mass is often thought to be the main factor regulating the metallicity of galaxies, we here show that the metallicity evolution of typical massive star-forming galaxies might in fact be the main driver of the observed molecular gas fraction, $M_{\text{mol}}/M_\star$, in galaxies through cosmic time.

5 CONCLUSIONS

We here presented a pilot survey targeting the CO emission counterparts of the host galaxies of strong H$_2$-absorbing GRBs at $z \sim 2$–3. We did not detect the redshifted CO(3–2) line in any of the three GRBs (GRBs 080607, 120815A, and 181020A). Assuming typical line ratios observed for star-forming galaxies at similar redshifts and a metallicity-dependent CO-to-H$_2$ conversion factor, we derived upper limits on the molecular gas masses of log ($M_{\text{mol}}/M_\star$) < 10.54, 11.67, 11.95, for the hosts of GRBs 080607, 120815A, and 181020A, respectively.

To place these H$_2$-selected GRB host galaxies into context, we compared them to the most recent compilation of GRB hosts with CO observations (Hatsukade et al. 2020, and references therein). First, the systems presented here expand the redshift range for which CO observations of GRB hosts have been obtained. Then, we examined them in terms of their molecular gas and stellar mass contents. While no strong conclusion can be inferred for the molecular-gas content of the hosts of GRBs 120815A and 181020A, we demonstrate that the host galaxy of GRB 080607 is globally deficient in molecules or ‘CO-dark’. This is surprising given the high metallicity ([X/H] $\geq -0.2$) and H$_2$ abundance ($N$(H$_2$) $= 10^{21.2}$ cm$^{-2}$) inferred from absorption, and the relatively high stellar mass inferred from emission (log ($M_\star/10^6$) = 10.45).

Motivated by the high inferred limits on the molecular gas masses of the hosts of GRBs 120815A and 181020A, resulting from their low absorption-derived metallicities and a metallicity-dependent CO-to-H$_2$ conversion factor $\alpha_{\text{CO}}$, we derived evolutionary tracks for $\alpha_{\text{CO}}$ as a function of redshift and stellar mass. We found that while the Galactic conversion factor $\alpha_{\text{CO, MW}}$ is suitable for massive galaxies with $M_\star \sim 10^{11}$ M$_\odot$ at $z \sim 0$–5, lower mass galaxies will show significantly higher $\alpha_{\text{CO}}$ values at all redshifts (by up to several orders of magnitudes). This will hamper the detection probability of CO in most star-forming galaxies at $z > 1$, since even large molecular gas reservoirs will show limited CO emission at these redshifts.

We demonstrated in the pilot survey presented here, resulting in non-detections of CO emission from GRB host galaxies, that the most feasible way to identify and study the molecular gas reservoirs in high-redshift, low-metallicity galaxies is through the detection of H$_2$ in absorption. Due to the metallicity-dependent CO-to-H$_2$ conversion factor $\alpha_{\text{CO}}$, these galaxies that otherwise show strong H$_2$ absorption will be too faint to be detected in emission using typical molecular gas tracers such as CO.

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DATA AVAILABILITY STATEMENT

The raw ALMA data are publicly available through the ALMA science archive. The reduced spectral cubes and source codes for the figures and tables presented in this manuscript are available from the corresponding author upon reasonable request.

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