The quantum interest conjecture in (3+1)-dimensional Minkowski space.

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The quantum inequalities, and the closely related quantum interest conjecture, impose restrictions on the distribution of the energy density measured by any time-like observer, potentially preventing the existence of exotic phenomena such as Alcubierre warp-drives or traversable wormholes. It has already been proved that both assertions can be reduced to statements concerning the existence or non-existence of bound states of a certain 1-dimensional quantum mechanical Hamiltonian. Using this approach, we will informally review a simple variational proof of one version of the Quantum Interest conjecture in (3+1)-dimensional Minkowski space.

Keywords: Quantum inequalities, quantum interest conjecture.

1. Introduction

Semiclassical general relativity predicts violations of the point-wise energy conditions associated with negative energy phenomena such as warp drives, traversable wormholes, and even time machines. However, the same theory can be used to constrain the magnitude and duration of negative energy pulses. Two of these restrictions are given by the quantum inequalities (QIs) and the quantum interest conjecture (QIC).

The QIs impose a lower bound on the expectation value (in a quantum state $\psi$) of the renormalized stress-energy tensor along a timelike geodesic,

$$I_{\psi,\omega} \equiv \oint \langle T_{00}^{\text{ren}}(t,0) \rangle_\psi \omega(t) \, dt,$$

weighted by a non-negative and normalized test function $\omega(t)$. The initial bounds on $I_{\psi,\omega}$ depended on the modified Bessel functions, and were obtained using a Lorentzian test function. Eventually Flanagan found a more general bound in (1+1)-dimensional Minkowski space for a massless scalar field,

$$I_{\psi,f} \geq -\frac{1}{16\pi} \oint \left( \frac{\omega(t)^2}{\omega(t)} \right)^2 \, dt.$$  (3)

Both inequalities, (2) and (3), can be written as a more general statement for $2m$-dimensional space-time,

$$\oint \langle T_{00}^{\text{ren}}(t,0) \rangle_\psi |f(t)|^2 \, dt \geq -\frac{1}{c_m} \oint |D^m f(t)|^2 \, dt,$$

with $\omega(t) = f^2(t)$. Here $D$ is the derivative operator, and the set of constants $c_m$ are given by

$$c_m = \left\{ \begin{array}{ll}
6\pi & m = 1; \\
\frac{m \pi m^{-1/2} 2^{2m} \Gamma(m - \frac{1}{2})}{m} & m \geq 2.
\end{array} \right. \quad (5)$$

The test function is now normalized such that $\int_{-\infty}^{+\infty} |f(x)|^2 \, dx = 1$. It is easy to check that we recover equations (2) and (3) from (4), by setting $m = 1$ and $m = 2$, respectively. Furthermore, by integrating by parts, the QIs become a statement regarding the lack of negative eigenvalues for a one-dimensional pseudo-Hamiltonian,

$$\langle f | H | f \rangle \geq 0,$$

where

$$H = (-1)^m D^{2m} + c_m V.$$  (7)

Here $V \equiv \langle T_{00}^{\text{ren}}(t,0) \rangle_\psi$ is effectively a potential for a quantum mechanical system. Therefore it is possible to use the point of view of one-dimensional quantum mechanics (after some technical considerations) to reduce the QIs to a much simpler framework.
2. The QIC as an eigenvalue problem

Once the quantum mechanical viewpoint is adopted, it is more convenient to also use its notation. Then the operator $H$ can be written as $P^2 m + V$, where $P$ and $V$ are operators in the usual Hilbert space of square-integrable functions. This technical construction allows us to rewrite the eigenvalue problem in coordinates, as an ordinary differential equation for the eigenfunctions $\varphi$ (the test functions $f$ used before); the multiharmonic time-independent Schrödinger equation (SDE),

\[ (-1)^{m} \frac{d^{2m}}{dx^{2m}} \varphi(x) + V(x) \varphi(x) = E \varphi(x), \]

where we have set $\hbar/2M \to 1$ to simplify the algebra. Again, if $m = 1$ we recover the (1+1)-dimensional case in the form of the time-independent SDE,

\[ -\frac{d^2}{dx^2} \varphi(x) + V(x) \varphi(x) = E \varphi(x). \]

Fortunately in this case, there is a theorem by Simon,\textsuperscript{10} which guarantees the existence of a negative eigenvalue for the SDE, if

\[ \int_{-\infty}^{\infty} V(x)\, dx \leq 0. \]

With this theorem, the QIC can be related to the QIs in (1+1) dimensions. It is also possible to reformulate the QIC for a more general set of energy pulses,\textsuperscript{11} unlike the original formulation, which is restricted to $\delta$-function pulses.\textsuperscript{12} To see this clearly, let us split the potential as

\[ V(x) = V(x)_+ - V(x)_-, \]

its positive part minus its negative part. Then in order to guarantee positive eigenvalues, the potential must violate (10). That is,

\[ \int |V(x)_+|\, dx > \int |V(x)_-|\, dx. \]

Since here the potential $V$ represents the expectation value of the renormalized stress-energy tensor, it is clear that to fulfill the QIs the net energy density must always be positive, i.e., its positive part must always overcompensate the negative one. This is a simplified version of the original formulation of the QIC.\textsuperscript{12}

3. The QIC in (3+1)-dimensional Minkowski space

The (3+1)-dimensional case can be recovered from the multiharmonic SDE by choosing $m = 2$. This yields the biharmonic SDE,

\[ \frac{d^4}{dx^4} \varphi(x) + V(x) \varphi(x) = E \varphi(x). \]

However, we now need to generalize Simon’s theorem for (13).

First, via a variational argument, the lowest eigenvalue $E$ of (13) satisfies

\[ E \leq \int [\varphi''(x)^2 + V(x) |\varphi(x)|^2] \, dx, \]

assuming all the test functions are normalized.

Secondly, let us choose the following test function

\[ \varphi_{test} = \sqrt{g(|x-\mu|/\sigma)}. \]

We then enforce the normalization

\[ \int g(|x|)\, dx = 1, \quad \int x g(|x|)\, dx = 0, \quad \int x^2 g(|x|)\, dx = 1, \]

to obtain

\[ \int |\varphi(x)|^2\, dx = 1, \quad \int x |\varphi(x)|^2\, dx = \mu, \]

and

\[ \int (x-\mu)^2 |\varphi(x)|^2\, dx = \sigma^2. \]

Note that the kinetic term of (14),

\[ \int \varphi''(x)^2\, dx = \frac{\int [g(|x|)]''^2\, dx}{\sigma^4}, \]

diverges. It contains a term proportional to $\delta$-function square, which arises from differentiating twice the absolute value. Nevertheless, by expanding the normalized functions $g(|x|)$ into a power series,

\[ g(|x|) = \sum_{n=0}^{\infty} a_n |x|^n, \]

we can get rid of the troublesome term by properly setting the expansion coefficient $a_1$ to zero. Moreover, to make (18) converge at zero, we also need $a_0 > 0$. The rest of the coefficients can be freely chosen. Then we have, from (14), (15) and (19)

\[ E\sigma^4 \leq \kappa + a_0 \sigma^3 \int V(x)\, dx + a_2 \sigma \int |x-\mu|^2 V(x)\, dx + a_3 \sigma \int |x-\mu|^3 V(x)\, dx + O(1/\sigma), \]
with \( \kappa = \int (|\sqrt{g}(x)|)^{2} dx \). Now choosing a sufficiently large \( \sigma \), it is clear that \( \int V(x) dx < 0 \) implies a negative eigenvalue for the differential equation.

Lastly, it is possible to collect more information from (20) if we set \( \int V(x) dx = 0 \). Then the next two terms become relevant. And since the sign and magnitude of \( a_{2} \) and \( a_{3} \) are arbitrary, either

\[
\forall \mu : \int |x - \mu|^{2} V(x) dx = 0, \tag{21}
\]

or

\[
\forall \mu : \int |x - \mu|^{3} V(x) dx = 0, \tag{22}
\]

is a sufficient condition to guarantee the absence of a bound state. Differentiating twice the last expression with respect to \( \mu \), we have

\[
\forall \mu : \int_{\mu}^{\infty} V(x) dx = 0, \tag{23}
\]

and finally \( V(x) = 0 \). That is, if \( \int V(x) dx = 0 \), a necessary condition for the lack of a bound state is that \( V(x) = 0 \). This proves the extension of Simon’s theorem for the biharmonic SDE, and it also proves the QIC in (3+1)-dimensional Minkowski space.

To clarify the proof of the QIC, we must recover the notation of semiclassical general relativity. Then the version of the QIC we just proved states that the QIs imply, either \( \langle T_{00}^{\text{ren}}(t, 0) \rangle_{\psi} = 0 \) everywhere along the world line, or

\[
\int \langle T_{00}^{\text{ren}}(t, 0) \rangle_{\psi} dt > 0, \tag{24}
\]

which is slightly stronger than the AWEC. Once again, splitting the energy density into its positive part minus its negative part,

\[
\int \langle T_{00}^{\text{ren}}(t, 0) \rangle_{+} dt > \int \langle T_{00}^{\text{ren}}(t, 0) \rangle_{-} dt, \tag{25}
\]

the positive energy density must overcompensate the negative part elsewhere along the world line.

4. Discussion

By proving the variant of Simon’s theorem for the biharmonic SDE, we were able to reformulate the QIC in (3+1)-dimensional flat spacetime similarly to the (1+1)-dimensional case studied in references 9 and 11. In flat spacetime, an energy pulse which satisfies the QIs (and the QIC proved above) must also fulfill an AWEC-like inequality.

Although several technical aspects have not been mentioned here, a more detailed proof of the QIC in flat spacetime can be found in reference 13.

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References

1. H. Epstein, V. Glaser and A. Jaffe, Nuovo Cim. 36 (1965) 1016; T. A. Roman, Phys. Rev. D 33, 3526 (1986).
2. M. Alcubierre, Class. Quant. Grav. 11 (1994) L73 [arXiv:gr-qc/0009013]; M. J. Pfenning and L. H. Ford, Class. Quant. Grav. 14 (1997) 1743 [arXiv:gr-qc/9702026]; F. S. N. Lobo and M. Visser, Class. Quant. Grav. 21, 5871 (2004) [arXiv:gr-qc/0406083]; F. S. N. Lobo and M. Visser, arXiv:gr-qc/0412065.
3. M. S. Morris and K. S. Thorne, Am. J. Phys. 56, 395 (1988); M. S. Morris, K. S. Thorne and U. Yurtsever, Phys. Rev. Lett. 61, 1446 (1988); M. Visser, Nucl. Phys. B 328, 203 (1989) [arXiv:0809.0927 [gr-qc]]; M. Visser, Phys. Rev. D 39, 3182 (1989) [arXiv:0809.0907 [gr-qc]]; M. Visser, S. Kar and N. Dadhich, Phys. Rev. Lett. 90, 201102 (2003) [arXiv:gr-qc/0301003].
4. J. Friedman, M. S. Morris, I. D. Novikov, F. Echeverria, G. Klinkhammer, K. S. Thorne and U. Yurtsever, Phys. Rev. D 42, 1915 (1990); S. W. Hawking, Phys. Rev. D 46, 603 (1992); M. Visser, arXiv:gr-qc/0204022.
5. L. H. Ford, Phys. Rev. D 43, 3972 (1991).
6. L. H. Ford and T. A. Roman, Phys. Rev. D 55, 2082 (1997) [arXiv:gr-qc/9607003].
7. E. E. Flanagan, Phys. Rev. D 56, 4922 (1997) [arXiv:gr-qc/9706006].
8. C. J. Fewster and S. P. Eveson, Phys. Rev. D 58, 084010 (1998) [arXiv:gr-qc/9805024].
9. C. J. Fewster and E. Teo, Phys. Rev. D 61, 084012 (2000) [arXiv:gr-qc/9908073].
10. B. Simon, Annals Phys. 97, 279 (1976).
11. E. Teo and K. F. Wong, Phys. Rev. D 66, 064007 (2002) [arXiv:gr-qc/0206066].
12. L. H. Ford and T. A. Roman, Phys. Rev. D 60, 104018 (1999) [arXiv:gr-qc/9901074].
13. G. Abreu and M. Visser, Phys. Rev. D 79, 065004 (2009) [arXiv:0808.1931 [gr-qc]].