Detailed study of the Milky Way globular cluster
Laevens 3.

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ABSTRACT

We present a photometric and spectroscopic study of the Milky Way satellite Laevens 3. Using MegaCam/CFHT g and i photometry and Keck II/DEIMOS multi-object spectroscopy, we refine the structural and stellar properties of the system. The Laevens 3 colour-magnitude diagram shows that it is quite metal-poor, old (13.0 ± 1.0 Gyr), and at a distance of 61.4 ± 1.0 kpc, partly based on two RR Lyrae stars. The system is faint (M_V = −2.8 ± 0.2 mag) and compact (r_h = 11.4 ± 1.0 pc). From the spectroscopy, we constrain the systemic metallicity ([Fe/H]_{spectro} = −1.8 ± 0.1 dex) but the metallicity and velocity dispersions are both unresolved. Using Gaia DR2, we infer a mean proper motion of (μ_α^* , μ_δ) = (0.51 ± 0.28, −0.83 ± 0.27) mas yr^{-1}, which, combined with the system’s radial velocity ⟨v_r⟩ = −70.2 ± 0.5 km s^{-1}, translates into a halo orbit with a pericenter and apocenter of 40.7 ± 5.6 and 85.6 ± 17.2 kpc, respectively. Overall, Laevens 3 shares the typical properties of the Milky Way’s outer halo globular clusters. Furthermore, we find that this system shows signs of mass-segregation which strengthens our conclusion that Laevens 3 is a globular cluster.

Key words: cluster: Globular – Local Group – object : Laevens 3

1 INTRODUCTION

In recent years, the faint regime of Milky Way (MW) satellites has been explored under the impulsion of large photometric surveys. Among those, we can cite the Sloan Digital Sky Survey (York et al. 2000), the Panoramic Survey Telescope and Rapid Response System 1 (Chambers et al. 2016) or the Dark Energy Survey (The Dark Energy Survey Collaboration 2005). These surveys led to numerous discoveries of faint satellites. Several old and metal-poor faint systems have been identified as globular clusters (GCs) (Balbinot et al. 2013, Laevens et al. 2014, Kim & Jerjen 2015, Kim et al. 2016), although some of them require confirmation (Martin et al. 2016c). Because of their old stellar populations, they can be considered as the witnesses of the formation of their host galaxy (Strader et al. 2005) and bring insights on low-mass galaxy formation. Furthermore, the chemodynamics of those GCs can also trace some of the current properties of their host (Pota et al. 2013). GCs can also be useful to constrain stellar population models (Chantereau, Charbonnel & Meynet 2016). The fact that these diffuse and small satellites survived for several billion years can also bring more information on their formation and internal processes (Baumgardt & Makino 2003, Renaud, Agertz & Gieles 2017).

The GCs associated with the MW span a wide range of luminosities, metallicities and distances (Harris 2010), but only a few have been discovered in the outer reaches of the halo (R_{gal} > 50 kpc). This specific group of clusters is in fact suspected to not have formed in-situ, but rather as companions in nearby dwarf galaxies and accreted at later times in the MW history (Mackey et al. 2010, Dotter, Sarajedini & Anderson 2011). While clusters like Pal 14 (Arp & van den Bergh 1960, d_{gal} ≈ 71 kpc) or AM-1 (Madore & Arp 1979, d_{gal} ≈ 125 kpc) have been known for decades, only a handful of fainter outer halo clusters were discovered in recent photometric surveys. Laevens 1/Crater (Belokurov et al. 2014, Laevens et al. 2014) and Kim 2 (Kim et al. 2015) fall in this
category. Such faint satellites often lie in the so-called “valley of ambiguity” where the frontier between dwarf galaxies and old stellar clusters is not clearly defined (Gilmore et al. 2007). Laevens 1 is a great illustration of that, as its very nature was disputed at the time of its discovery. Indeed, while Laevens et al. (2015) identified the system as a cluster, Belokurov et al. (2014) proposed that the satellite could have been a tidally disrupted dwarf galaxy. This example only accentuates the hardness of studying these faint, distant stellar systems. In such an extreme regime, the combination of photometric, chemical, and kinematics data is needed to both classify and understand those systems.

Laevens 3 (Lae 3) is a system first discovered in the Pan-STARRS 1 (Chambers et al. 2016) PS1 data by Laevens et al. (2015). At the time, it was found to be compact (FH = 7 ± 2 pc) and the existence of an RR Lyrae star in this region, probably belonging to the system, allowed to constrain the distance to the system (64 ± 3 kpc). Using this distance, Laevens et al. (2015) found that the main sequence of Lae 3 was compatible with a stellar population of 8 Gyr, and a metallicity of [Fe/H] = −1.9. From these properties, the authors concluded that the system is a faint Milky Way globular cluster.

In this work, we undertake a careful refinement of the properties of the satellite through deep broadband photometry with CFHT/MegaCam, as well as the first spectroscopic follow-up of the system using Keck/DEIMOS (Faber et al. 2003). Section 2 discusses the technical aspects of our observations. Section 3 details the photometric analysis that derives the structural and colour-magnitude diagram (CMD) properties of the satellite. In section 4, we present the dynamics of Lae 3 using multi-object spectroscopy, while section 5 details the orbital properties of the satellite obtained with the Gaia Data Release 2 data. Finally, the nature and main properties of Lae 3 are discussed in section 6.

2 OBSERVATIONS

2.1 Photometry

The photometry used in this work consists of multi-exposures MegaCam broadband g- and i-band images. The exposure times are of 3 × 480 s for g and 3 × 540 s for i. The observations were conducted in service mode by the CFHT crew during the night of July 18th, 2015 under excellent seeing conditions (≈ 0.3″), and the data reduced following the procedure detailed in Longeard et al. (2018 L18). We use the Cambridge Astronomical Survey Unit (CASU, Irwin & Lewis 2001) pipeline flags to perform the star/galaxy separation. CASU also indicates all saturated sources. The calibration of the MegaCam photometry (Boulade et al. 2003) is performed onto the PS1 photometric system similarly to L18. We first cross-identified all unsaturated point sources between PS1 and MegaCam. Only stars with photometric uncertainties below 0.05 in both catalogs are then considered for the calibration. We assume that the transformation between the PS1 and MegaCam photometry can be reliably modelled by a second-order polynomial, with a 3-σ clipping procedure.

All stars saturated in the MegaCam photometry are directly imported from the PS1 catalog, for a total of 51,759 stars. Finally, the catalog is dereddened using the 2D dust map from Schlegel, Finkbeiner & Davis (1998) to determine the line-of-sight extinction and Schlafly & Finkbeiner (2011) for the extinction coefficients.

2.2 Spectroscopy

The spectroscopic run for Lae 3 was performed on the night of Sept 7, 2015 (Julian date of 2457272.5) using Keck II/DEIMOS. The targets were selected based on their distance to Lae 3 and their location on the colour-magnitude diagram, using the PS1 photometry presented in Laevens et al. (2015). A total of 51 stars were observed using the OG550 filter and the 1200 lines mm⁻¹ grating. The typical central wavelength resolution is R ≈ 8500, covering the spectral range from 6500 to 9000 Å. The spectra were then processed using the IRAF SIMULATOR package from the Keck Observatories and the pipeline detailed in Ibata et al. (2011). Stars with a signal-to-noise ratio below 3 as well as the ones with radial velocity uncertainties above 15 km s⁻¹ are discarded from the spectroscopic catalog. The resulting catalog consists of 44 stars for which the spatial and CMD distributions are shown in Figure 1. Finally, the instrumental systematic velocity uncertainty is chosen to be the same as in Longeard et al. (2019), with δ thr = 1.8 ± 0.2 km s⁻¹.

3 BROADBAND PHOTOMETRY ANALYSIS

The region including Lae 3 is shown in the left panel of Figure 1 with the stars observed spectroscopically colour-coded by their velocities. The central region of the system is densely populated. The colour-magnitude diagram within two half-light radii of Lae 3 is shown in the right panel of Figure 1. The great depth of the MegaCam photometry allows us to probe the system two magnitudes below the Main Sequence Turn-Off (MSTO) and clearly reveals the main sequence of Lae 3. Our spectroscopic sample extends all the way down to the sub-giant branch, and suggests that Lae 3 possesses at least a few Red Giant Branch (RGB) stars. Four RR Lyrae stars are located in the vicinity of the satellite according to the catalog of Sesar et al. (2017). Among those, only stars with a RRab classification score greater than 90 per cent are selected, as the distance modulus measurement of RRc stars can be biased. Two stars pass this criterion and have a m – M of 18.87 ± 0.06 and 18.89 ± 0.06 mag respectively. By doing the mean of these two distance moduli, we obtain a distance modulus estimate of 18.88 ± 0.04 mag for Lae 3 (59.7 ± 0.2 kpc in physical distance).

3.1 Structural and CMD fitting

We aim to derive the structural and stellar population properties of Lae 3. As such, we rely on the technique presented in Martin et al. (2016b) and L18. The stellar population parameters that we aim to infer are the age A, metallicity [Fe/H]CMD, the α abundance ratio [α/Fe], and the distance modulus m – M. The structural properties that are determined are the spatial offsets of the centroid from the literature values (α = +316.72635°, δ = +14.98000°) X₀ and
Figure 1. Left panel: Spatial distribution of the Lae 3-like stellar population in the field of view. The CFHT image of the 2.5’×2.5’ region around Lae 3 in the i band is shown in the upper-left corner. The red circle represents the two half-light radii ($r_h \sim 0.64'$) region of Lae 3. The two RR Lyrae identified in the system are shown as magenta stars. The spectroscopic dataset is represented by circles, colour-coded according to their heliocentric velocities. Filled circles stand for stars identified as Lae 3 members. The two RR Lyrae identified in the system are shown as magenta stars. The spectroscopic dataset is represented by circles, colour-coded according to their heliocentric velocities. Filled circles stand for stars identified as Lae 3 members. Right panel: CMD within two half-light radii of Lae 3. The best fitting isochrone derived in section 3.1 is represented as a solid green line, while the stellar population inferred without any distance or metallicity priors is represented by the light green dashed line. Photometric uncertainties are reported as grey error bars on the left side of the plot.

$Y_0$, the ellipticity ε, the half-light radius $r_h$, the position angle of the major axis $\theta$, and the number of stars $N_\ast$ of the system within the dataset.

To derive the structural parameters, the satellite is assumed to follow an exponential radial density profile, while the spatial density of the background is assumed to be constant over the field. The stellar characteristics are determined by assuming that the CMD of the satellite can be considered as the sum of two components: a unique stellar population for Lae 3, and a contamination from the foreground stars. Given the appearance of the Lae 3 sequence in Figure 1, these assumptions are reasonable as the differences between isochrones in the metal-poor regime are not significant, except in the case of important spreads in both age and metallicity. The modelling of the CMD contamination is done empirically, by selecting all stars outside 5$r_h$ of the system. The CMD of this subsample is further binned and smoothed by a gaussian kernel of 0.1 in both colour and magnitude. The Lae 3 stellar population is, on the other hand, modelled using old and metal-poor isochrones from the Darmouth library [Dotter et al. 2008]. The Lae 3 likelihood model is built by convolving each isochrone track by the typical photometric uncertainties of the data at a given

$$(g_0, i_0)$$

This model is then weighted by both the luminosity function of the track considered, and the completeness of the data at a given $(g_0, i_0)$. This method is discussed in further details in L18.

The distance inferred using the RR Lyrae in the field can be used as a prior for our analysis. Moreover, and anticipating on section 4, the spectroscopic analysis of three bright Lae 3 member stars allows us to infer the metallicity of the satellite to be $< [\text{Fe/H}]_{\text{spectro}} >= -1.8 \pm 0.1$ dex. The PDF of this result can also be used as a prior.

The structural and CMD parameters are inferred all together and the results are displayed in Table 1, while the PDFs are shown in figure 2. We find that Lae 3 is spherical, with a half-light radius of 0.64 ± 0.05 arcminutes that translates into a physical $r_h$ of 11.4 ± 1.0 pc. The measured half-light radius is larger than that of the discovery paper [Laevens et al. 2015 ~ 0.4']. To investigate this discrepancy, the sample is split between bright (15.0 < $g_0$ < 23.5) and faint (24.0 < $g_0$ < 25.0) stars, and the structural properties of Lae 3 are derived in both cases. A significant difference arises in terms of half-light radius as shown in Figure 3. The sample of bright stars yields a more compact size than with the faint-end of the population. Such a discrepancy would naturally arise in a satellite in which a mass-segregation process has already occurred, and could explain the difference between this work and [Laevens et al. 2015] who analysed...
the system with the shallower PS1 data. To test this, the structural analysis is performed using directly the PS1 data. The resulting PDF is shown as the dashed line in Figure 3. The half-light radius inferred with this procedure is similar to the one obtained by L15, suggesting that the larger size derived from the MegaCam data is driven by less massive stars below $g < 22.5$ mag and that Lae 3 is mass-segregated. We compute the relaxation time of Lae 3 using the equations of Koposov et al. (2007) and references therein to confirm that the satellite had enough time to mass-segregate. We choose a mass-to-light ratio of 2 expected from old GCs (Bell & de Jong 2001), a total luminosity of $1125\ L_\odot$ determined below, and an average star mass of $0.6\ M_\odot$. The resulting half-light relaxation time is around 2.2 Gyr, largely smaller than our inference of the age of the satellite ($13.0 \pm 1.0$ Gyr).

Two favoured stellar populations are presented in Figure 2 with and without using the priors on the metallicity and distance modulus coming respectively from the spectro-
scopistic analysis of section 4 and the two RR Lyrae in the system. Without those priors, Lae 3 is found to be old (13.0±1.0 Gyr) and metal-poor (< [Fe/H]_CMD > = −2.0 ± 0.1 dex). The abundance ratio in α elements is [α/Fe] = 0.2 ± 0.2 dex, while the distance modulus is m−M = 19.05±0.02 mag, i.e. a physical distance of 64.4±3.6 kpc. This model is represented as a dashed light green line in Figure 4 and nicely follows the sequence of the satellite and the spectroscopic members identified in the next section. The favoured model, i.e. the one based on the metallicity and distance priors, is similar. The structural properties, age, metallicity and α abundance ratio are compatible. However, the satellite is found to be closer (m−M = 18.94±0.05 mag, which translates in a physical distance of 61.4±1.2 kpc) in this case. This population, represented as a solid green line in Figure 4, also follows the features of Lae 3 in the CMD. The two isochrones are barely distinguishable and the last model is the one used in the rest of this work since it is based on a spectroscopic measurement of the metallicity of the system. Using the favoured model, two quantities are defined: a “CMD probability membership” that assigns a probability to a given star solely based on its compatibility with the favoured stellar population of Lae 3 and a “CMD and spatial probability membership” that also takes the spatial location of a given star into account.

Using this CMD membership probability, we search for potential tidal structures. To do so, the field of view is spatially binned with 0.2 arcminutes bins. The CMD probability of all stars falling in a given bin are then added. This procedure therefore assigns higher values to bins that contain stars compatible with the stellar population of Lae 3. The result is shown in Figure 4. This analysis shows that the satellite is highly spherical and that there is no tidal feature in the field of view compatible with the CMD properties of Lae 3.

The luminosity of the satellite is estimated following the method detailed in [Martin et al. (2016a)] which consists in simulating thousands of CMDs with the stellar and structural properties of Lae 3 derived earlier, and compute their resulting luminosities. This procedure yields a luminosity of L_V = 1125^{+221}_{−129} L_⊙, translating into an absolute magnitude of M_V = −2.8^{+0.2}_{−0.3} mag. This result is roughly one magnitude fainter than that found by [Laevens et al. (2015)] in the discovery paper of Lae 3. We observed a similar trend for another faint satellite discovered by [Laevens et al. (2015)]: Draco II (Dra II). In L18, the inferred luminosity was significantly lower than found in the 2015 paper, and we concluded that it is most likely due to the overestimation of the number of giants in the system, probably due to the shallowness of the PS1 data used for the discovery of both Lae 3 and Dra II. Though Lae 3 is clearly brighter than Dra II, it is also significantly more distant, and the same overestimation effect might have affected the result of [Laevens et al. (2015)], as using the same technique for a brighter MW satellite ([Longeard et al. 2019]) did not yield such an effect.


4 SPECTROSCOPIC ANALYSIS

The distribution of the heliocentric velocities for all stars in our spectroscopic sample is shown in the top panel of Figure 4 along with their radial distances and spectroscopic metallicities (if possible).

4.1 Dynamical properties

The Lae 3 population is not prominent, and its systemic velocity overlaps that of the foreground MW stars (Figure 4). Our approach is similar to L18: the velocity distribution is assumed to be the sum of the contamination (halo and disc stars) and the Lae 3 population, both modelled with different normal distributions. To highlight Lae 3’s population in the spectroscopic dataset, the individual likelihood of each star is weighted by its spatial and CMD probability estimated from the favoured structural model of section 3. This analysis yields a systemic radial velocity of $v_r = -70.2 \pm 0.5 \text{ km s}^{-1}$. The 1-D marginalised PDFs of the velocity parameters are represented in the left panels of Figure 6. As a consequence to the low number of Lae 3 stars, the velocity dispersion is unresolved. Finally, six stars with a dynamical, structural and CMD membership probability greater than 90 per cent are identified as Lae 3 members and shown as filled circles in Figure 4.

4.2 Spectroscopic metallicity

The individual metallicities of stars observed with spectroscopy can be estimated using the calibration of the Calcium triplet for RGB stars, and shown in Figure 5. Member stars fainter than 21 in the $g$ band, and with $S/N < 10$ are further discarded from our spectroscopic catalog. Only three stars are left to infer the systemic metallicity and metallicity dispersion of Lae 3. By assuming that the metallicities are normally distributed, this yields a spectroscopic metallicity of $\langle \text{[Fe/H]} \rangle_{\text{spectro}} = -1.8 \pm 0.1 \text{ dex}$. The same analysis is also performed using the calibration for metal-poor stars on the RGB and sub-RGB branch, and yields compatible results. Once more, low number statistics has a direct consequence on our ability to constrain efficiently the metallicity dispersion, which is found to be unresolved, with $\sigma_{\text{[Fe/H]}} < 0.5 \text{ dex}$ at the 95 per cent confidence level. The PDFs of both parameters are shown in the right panels of Figure 6.

| Parameter | Unit | Prior | Favoured model | Uncertainties |
|-----------|------|-------|----------------|--------------|
| RA $\alpha$ | degrees | 316.72938021 | ±0.00076375 |
| DEC $\delta$ | degrees | 21:06:55:05 | ±0.001 |
| $l$ | degrees | 63.598 | ±0.001 |
| $b$ | degrees | -21.176 | ±0.001 |
| $r_h$ | arcmin | 0.64 | ±0.05 |
| $r_h$ | pc | > 0 | ±1.0 |
| $\theta$ | degrees | 72 | ±24 |
| $e$ | — | > 0 | ±0.11 |
| Distance modulus | mag | $G(18.88, 0.04)$ | ±0.09 |
| Distance | kpc | 61.4 | ±1.2 |
| Age | Gyr | [8.0, 13.5] | ±0.92 |
| $\langle \text{[Fe/H]} \rangle_{\text{spectro}}$ | dex | — | ±0.1 |
| $[\alpha/\text{Fe}]$ | dex | [-0.2, 0.6] | ±0.2 |
| $M_V$ | mag | — | ±0.2 |
| $\mu_0$ | mag arcsec$^{-2}$ | 25.0 | ±0.3 |
| $< v_r >$ | km s$^{-1}$ | — | ±0.5 |
| $\mu$ | mas yr$^{-1}$ | 0.51 | ±0.28 |
| $\mu_s$ | mas yr$^{-1}$ | — | ±0.27 |
| Apocenter | kpc | 85.6 | ±17.2 |
| Pericenter | kpc | 40.7 | ±5.6 |
| $\epsilon_{\text{orbit}}$ | — | > 0 | ±0.04 |
| $U$ | km s$^{-1}$ | 13.1 | ±6.4 |
| $V$ | km s$^{-1}$ | — | ±5.6 |
| $W$ | km s$^{-1}$ | — | ±5.6 |
| $L_z$ | km s$^{-1}$ kpc | 793 | ±4010 |
| $E$ | km$^2$ s$^{-2}$ | 20819 | ±1482 |

Table 1. Inferred properties of Lae 3.
LAEVENS 3 COMES OUT AS A FAIRLY METAL-POOR STELLAR SYSTEM:

- [Fe/H] = −0.83 ± 0.27 dex which places Lae 3 far off the luminosity-metallicity relation of DGs (Kirby et al. 2013) as shown in Figure 9. The metallicity dispersion is also unresolved. Similarly to Laevens et al. (2015), two RR Lyrae stars are used to estimate the distance of Lae 3, and yield a distance modulus of 18.88 ± 0.04 mag. Using these results as priors, we derive the structural and CMD properties of the satellite. Lae 3 shows the main characteristics of MW globular clusters: the satellite is fairly spherical and is at the same luminosity (McConnachie 2012; Kirby et al. 2013), such as Ret II (M_V ≈ −2.7), Hor I (M_V ≈ −3.4) or Boo II (M_V ≈ −2.7) as shown in the bottom-left panel of Figure 9. Regarding the size and galactocentric distance of the satellite, Lae 3 can be compared to SMASH 1 (Martin et al. 2016). SMASH 1 has a size of 9.1^{+3.9}_{-3.4} pc and is lying at ~57 kpc of the center of the galaxy. The location and distance of SMASH 1 imply that it may be a satellite of the LMC. However, Lae 3 is brighter (~2.8 vs ~1.0 mag) and is more metal-rich (~1.8 vs ~2.2 dex). The top-left panel of Figure 9 shows that the systemic metallicity of Lae 3 is offset by ~0.7 dex from the metallicity-luminosity relation of dwarf galaxies (Kirby et al. 2013). We have to turn to Pal 1 or Pal 13 (Harris 2010) to find a cluster with a luminosity comparable to the one of Lae 3 (respectively of ~2.5 and ~3.8 mag). Still, these two GCs are much more compact, with a size of the order of the parsec.

Both the velocity and metallicity dispersions of Lae 3 are unresolved, although the small number of member stars in our spectroscopic dataset does not give stringent enough

5 GAIA DR2 PROPER MOTIONS AND ORBIT

To infer the orbital properties of Lae 3, we cross-match all spectroscopic members and RR Lyrae stars with the Gaia Data Release 2 (Gaia Collaboration et al. 2018). Among those, four stars have a proper motion (PM) measurement in Gaia. Furthermore, all stars in the Gaia catalog with a CMD and structural membership probability greater than 90 per cent are included. Six additional stars are retrieved through this procedure, and their PMs are compatible with those of the spectroscopic members, as shown in Figure 7. The uncertainty-weighted average PM of Lae 3 yields μ_α Lae3 = μ_δ Lae3 cos(δ) = 0.51 ± 0.28 mas yr^{-1} and μ_δ Lae3 = −0.83 ± 0.27 mas yr^{-1}. These measurements take into account the systematic error of 0.035 mas yr^{-1} on the PMs for dSph as shown by Helmi et al. (2018). We point out that this choice of systematic error does not change our results, given the measured uncertainties on the PM of the satellite.

We use the GALPY package (Bovy 2015) to integrate the orbit of Lae 3. The MW potential chosen to integrate Lae 3 orbit is a variant of the “MWPotential14” defined within GALPY, but updated with a halo mass of 1.2x10^{12} M_☉ (Bland-Hawthorn & Gerhard 2016). Five thousand orbits are integrated backwards and forwards over 5 Gyr, each time by randomly drawing a position, distance, radial velocity, and PMs from their corresponding PDFs. Around 20 per cent of the resulting orbits are not bound to the MW. In the case where Lae 3 is bound to the MW, the pericenter is at 40.7^{+5.6}_{-14.7} kpc and the apocenter is at 85.6^{+17.2}_{-7} kpc. The favoured orbit of the satellite is shown as a solid blue line in Figure 8 and corresponds to a typi-
Figure 6. Left panels: 1-D marginalised PDFs of the systemic velocity and its associated dispersion. Right panels: 1-D marginalised PDFs of the systemic metallicity and its associated dispersion. The two measurements of the dispersions are unresolved.

Figure 7. PMs of all stars within 15′ of Lae 3. The grey transparent dots show the PMs of field stars. The measurements of the four spectroscopic members with PM in Gaia DR2 are represented as squares, while the red stars and dots respectively show the PMs of the RR Lyrae stars as well as the spatially and CMD selected stars. The large green dot marks the combined PM measurement of Lae 3.

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Figure 8. Orbits of Lae 3 in the X-Y, X-Z and Y-Z planes integrated over 5 Gyr. The blue line is the orbit for the favoured distance, radial velocity, position and proper motion. Grey, transparent lines are random realizations of the orbit. The MW is represented by the black circle ($R_{MW} = 15$ kpc), while the blue dot indicates the location of Lae 3 at present day.

to have the opportunity to conduct observations from this mountain.

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REFERENCES

Arp H., van den Bergh S., 1960, PASP, 72
Balbinot E. et al., 2013, ApJ, 767, 101
Baumgardt H., Makino J., 2003, MNRAS, 340, 227
Bechtol K. et al., 2015, ApJ, 807, 50
Bell E. F., de Jong R. S., 2001, ApJ, 550, 212
Belokurov V., Irwin M. J., Koposov S. E., Evans N. W., Gonzalez-Solares E., Metcalfe N., Shanks T., 2014, MNRAS, 441, 2124
Belokurov V. et al., 2010, ApJ, 712, L103
—, 2009, MNRAS, 397, 1748
—, 2007, ApJ, 654, 897
—, 2006, ApJ, 647, L111
Bland-Hawthorn J., Gerhard O., 2016, ARAA, 54, 529
Boulade O. et al., 2003, 4841, 72
Bovy J., 2015, ApJSupp, 216, 29
Carrera R., Pancino E., Gallart C., del Pino A., 2013, MNRAS, 434, 1681

Carretta E., Bragaglia A., Gratton R. G., Leone F., Recio-Blanco A., Lucatello S., 2006, A&A, 450, 523
Carretta E. et al., 2009, A&A, 505, 117
—, 2007, A&A, 464, 967
Carretta E., Lucatello S., Gratton R. G., Bragaglia A., D’Orazi V., 2011, A&A, 533, A69
Chambers K. C. et al., 2016, arXiv:1612.05560
Chantereau W., Charbonnel C., Meynet G., 2016, A&A, 592, A111
Cohen J. G., Kirby E. N., Simon J. D., Geha M., 2010, ApJ, 725, 288
Collins M. L. M. et al., 2010, MNRAS, 407, 2411
Conn B. C., Jerjen H., Kim D., Schirmer M., 2018, ApJ, 852, 68
Dotter A., Chaboyer B., Jevremović D., Kostov V., Baron E., Ferguson J. W., 2008, ApJSupp, 178, 89
Dotter A., Sarajedini A., Anderson J., 2011, ApJ, 738, 74
Drlica-Wagner A. et al., 2015, ApJ, 813, 109
Faber S. M. et al., 2003, 4841, 1657
Gaia Collaboration, Brown A. G. A., Vallenari A., Prusti T., de Bruijne J. H. J., Babusiaux C., Bailer-Jones C. A. L., 2018, arXiv:1804.09365
Figure 9. Comparison of Lae 3 with other GCs and dwarf galaxies of the Milky Way. Squares represent dwarf galaxies while circles represent globular clusters, and the diamond corresponds to Lae 3. Triangles stand for recently discovered dwarf-galaxy candidates that await confirmation. Hollow markers correspond to systems for which no metallicity dispersion measurement can be found in the literature. The solid line in the top-left panel corresponds to the luminosity-metallicity relation of Kirby et al. (2013) for dwarf spheroidals and dwarf irregulars. Dashed lines represent the RMS about this relation, also taken from Kirby et al. (2013). Among the 123 globular clusters presented here, the properties of 116 were extracted from Harris (1996) catalog, revised in 2010. For the remaining ones (Kim 1, Kim 2, Kim 3, Laevens 1, Balbinot 1, Munoz 1 and SMASH 1) parameters of the discovery publications were used (Kim & Jerjen (2015), Kim et al. (2015), Kim et al. (2016), Laevens et al. (2014), Balbinot et al. (2013), Munoz et al. (2012) and Martin et al. (2016c)). Globular cluster metallicity spread measurements are taken from Willman & Strader (2012) and references therein; Carretta et al. (2006, 2007, 2009, 2011), Cohen et al. (2010), Gratton et al. (2007), Johnson & Pilachowski (2010), and Marino et al. (2011), McConnachie (2012) and Willman & Strader (2012) are used to compile the properties of the dwarf galaxies represented here. The 18 dwarf galaxies represented here are: Bootes 1 (Belokurov et al. 2006), Norris et al. (2010), Canes Venatici I (Zucker et al. 2006b), Canes Venatici II (Sakamoto & Hasegawa 2000), Coma Berinices, Hercules, Leo IV and Segue 1 (Belokurov et al. 2007), Draco and Ursa Minor (Wilson 1955), Fornax (Shapley 1938), Leo I and Leo II (Harrington & Wilson 1950), Pisces II (Belokurov et al. 2010), Sculptor (Shapley 1938a), Sextans (Irwin et al. 1996), Ursa Major I (Willman et al. 2005a), Ursa Major II (Zucker et al. 2006a), Willman I (Willman et al. 2005a). Their metallicity and metallicity spreads were drawn from Kirby et al. (2008), Kirby et al. (2010), Norris et al. (2010), Willman et al. (2011). The dwarf galaxy candidates discovered recently and shown on this figure are Bootes II (Koch & Rich 2014), DESI (Luque et al. 2016), Eridanus III (Bechtol et al. 2015), Conn et al. 2018, Koposov et al. 2015), Hyades II (Martin et al. 2015), Pegasus III (Kim & Jerjen 2015), Retiuchum II and Horologium I (Koposov et al. 2015a), Segue II (Belokurov et al. 2009), and the most significant candidates of Drlica-Wagner et al. (2015): Gru II, Tuc III, and Tuc IV.
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Table 2. Properties of our spectroscopic sample. Confirmed members are denoted by “Y” and non-members by “N”. The star denoted “Y?” is a plausible member, as its position, spectroscopic metallicity and velocity are compatible with those of Lae 3. However, it is not confirmed as it does not pass our membership probability cut ($P_{\text{mem}} \geq 90\%$).
Gilmore G., Wilkinson M. I., Wyse R. F. G., Kleyna J. T., Koch A., Evans N. W., Grebel E. K., 2007, ApJ, 663, 948
Gratton R. G. et al., 2007, A&A, 464, 933
Harrington R. G., Wilson A. G., 1950, PASP, 62, 118
Harris W. E., 1996, AJ, 112, 1487
—, 2010, ArXiv e-prints
Helmi A., Babusiaux C., Koppelman H. H., Massari D., Veljanoski J., Brown A. G. A., 2018, ArXiv e-prints
Ibata R., Sollima A., Nipoti C., Bellazzini M., Chapman S. C., Dalessandro E., 2011, ApJ, 738, 186
Irwin M., Lewis J., 2001, New Astronomy Review, 45, 105
Irwin M. J., Buncclark P. S., Bridgeland M. T., McMahon R. G., 2001, MNRAS, 320, 661
Johnson C. I., Pilachowski C. A., 2010, ApJ, 722, 1373
Kim D., Jerjen H., 2015, ApJ, 799, 73
Kim D., Jerjen H., Mackey D., Da Costa G. S., Milone A. P., 2016, ApJ, 820, 119
Kim D., Jerjen H., Milone A. P., Mackey D., Da Costa G. S., 2015, ApJ, 803, 63
Kirby E. N., Cohen J. G., Guhathakurta P., Cheng L., Bullock J. S., Gallazzi A., 2013, ApJ, 779, 102
Kirby E. N. et al., 2010, ApJSupp, 191, 352
Kirby E. N., Simon J. D., Geha M., Guhathakurta P., Frebel A., 2008, ApJ, 685, L43
Koch A., Rich R. M., 2014, ApJ, 794, 89
Koposov S. et al., 2007, ApJ, 669, 337
Koposov S. E., Belokurov V., Torrealba G., Evans N. W., 2015a, ApJ, 805, 130
Koposov S. E. et al., 2015b, ApJ, 811, 62
Laevens B. P. M. et al., 2015, ApJ, 813, 44
—, 2014, ApJ, 786, L3
Longeard N. et al., 2018, MNRAS, 480, 2609
—, 2019, arXiv:1902.0780
Luque E. et al., 2016, MNRAS, 458, 603
Mackey A. D. et al., 2010, ApJ, 717, L11
Madore B. F., Arp H. C., 1979, ApJ, 227, L103
Marino A. F. et al., 2011, A&A, 532, A8
Martin N. F. et al., 2016a, MNRAS, 458, L59
—, 2016b, ApJ, 833, 167
—, 2016c, ApJ, 830, L10
—, 2015, ApJ, 804, L5
McConnachie A. W., 2012, AJ, 144, 4
Muñoz R. R., Geha M., Côté P., Vargas L. C., Santana F. A., Stetson P., Simon J. D., Djorgovski S. G., 2012, ApJ, 753, L15
Norris J. E., Wyse R. F. G., Gilmore G., Yong D., Frebel A., Wilkinson M. I., Belokurov V., Zucker D. B., 2010, ApJ, 723, 1632
Pota V. et al., 2013, MNRAS, 428, 389
Renaud F., Agertz O., Gieles M., 2017, MNRAS, 465, 3622
Sakamoto T., Hasegawa T., 2006, ApJ, 653, L29
Schlafly E. F., Finkbeiner D. P., 2011, ApJ, 737, 103
Schlegel D. J., Finkbeiner D. P., Davis M., 1998, ApJ, 500, 525
Sesar B. et al., 2017, AJ, 153, 204
Shapley H., 1938a, Harvard College Observatory Bulletin, 908, 1
—, 1938b, Nature, 142, 715
Starkenburg E. et al., 2010, A&A, 513, A34
Strader J., Brodie J. P., Cenarro A. J., Beasley M. A., Forbes D. A., 2005, AJ, 130, 1315
The Dark Energy Survey Collaboration, 2005, arXiv:astro-ph/0510346
Willman B. et al., 2005a, AJ, 129, 2692
—, 2005b, ApJ, 626, L85
Willman B., Geha M., Strader J., Strigari L. E., Simon J. D., Kirby E., Ho N., Warren A., 2011, AJ, 142, 128
Willman B., Strader J., 2012, AJ, 144, 76
Wilson A. G., 1955, PASP, 67, 27
York D. G. et al., 2000, The Astronomical Journal, 120, 1579
Zucker D. B. et al., 2006a, ApJ, 650, L41
—, 2006b, ApJ, 643, L103