Orbital phase resolved spectroscopy of
4U1538–52 with MAXI

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ABSTRACT

Context. 4U 1538–52, an absorbed high mass X-ray binary with an orbital period of ~3.73 days, shows moderate orbital intensity modulations with a low level of counts during the eclipse. Several models have been proposed to explain the accretion at different orbital phases by a spherically symmetric stellar wind from the companion.

Aims. The aim of this work is to study both the light curve and orbital phase spectroscopy of this source in the long term. Particularly, the folded light curve and the changes of the spectral parameters with orbital phase to analyse the stellar wind of QV Nor, the mass donor of this binary system.

Methods. We used all the observations made from the Gas Slit Camera on board MAXI of 4U 1538–52 covering many orbits continuously. We obtained the good interval times for every orbital phase range which were the input to extract our data. We estimated the orbital period of the system

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and then folded the light curves and we fitted the X-ray spectra with the same model for every orbital phase spectrum. We also extracted the averaged spectrum of all the MAXI data available. 

Results. The MAXI spectra in the 2–20 keV energy range were fitted with an absorbed Comptonization of cool photons on hot electrons. We found a strong orbital dependence of the absorption column density but neither the fluorescence iron emission line nor low energy excess were needed to fit the MAXI spectra. The variation of the spectral parameters over the binary orbit were used to examine the mode of accretion onto the neutron star in 4U 1538–52. We deduce a best value of $\dot{M}/v_{\infty} = 0.65 \times 10^{-9} M_\odot \text{yr}^{-1}/(\text{km s}^{-1})$ for QV Nor.

Key words. X-rays: binaries – pulsars: individual: 4U 1538–52

1. Introduction

4U 1538–52 is a high mass X-ray binary (HMXB) pulsar with a B-type supergiant companion, QV Nor. The orbital period of the binary system is $\sim 3.73$ days (Davison et al. (1977); Clark (2000); Mukherjee et al. (2006)), and the magnetised neutron star has a spin period of $\sim 529$ s (Davison et al. (1977); Becker et al. (1977)). The distance to the source was estimated to be between 4.0 and 7.4 kpc (Crampton et al. (1978); Ilovaisky et al. (1979); Reynolds et al. (1992)), and we use its latest estimation in this work, i.e., 6.4 kpc. Although an orbital low eccentricity was estimated to be $e \sim 0.08$, (Corbet et al. (1993)), a higher value of $e \sim 0.17$ suggested an elliptical orbit for this system (Clark (2000); Mukherjee et al. (2006)). Nevertheless, Rawls et al. (2011) argued that a circular orbit and an eccentric orbit cannot be distinguished. The X-ray luminosity is minimum during the eclipse which lasts $\sim 0.6$ days (Becker et al. (1977)).

Both timing and spectral properties of 4U 1538–52 have been studied with some X-ray observatories in different energy bands, such as Tenma (Makishima et al. (1987)), EXOSAT (Robba et al. (1992)), Ginga (Bulik et al. 1993), RXTE (Clark (2000)), BeppoSAX (Robba et al. (2001)), XMM-Newton (Rodes-Roca et al. (2011)), INTEGRAL (Hemphill et al. 2013), and Suzaku (Hemphill et al. 2014). The X-ray characteristics derived from these analyses suggest that the mass accretion onto the neutron star is consistent with a spherically symmetric stellar wind from the companion star (Mukherjee et al. 2006). The X-ray spectrum based on a broad band energy range has been described by different absorbed power-law relations, modified by a high-energy cutoff including a fluorescence iron line at 6.4 keV, whenever present, and with cyclotron resonant scattering features at $\sim 21$ keV (Clark et al. (1990); Robba et al. (2001)) and at $\sim 47$ keV (Rodes-Roca et al. (2009)). A soft excess component is found in the X-ray spectrum from BeppoSAX and XMM-Newton observations (Robba et al. (2001); Rodes-Roca et al. (2011)). Similar to other wind-fed pulsars, the X-ray flux changes while the neutron star is moving in its orbit around the circumstellar environment of the companion star. Therefore, the variability is associated with the mass-loss rate by stellar wind captured by the gravitational field of the neutron star ($M$). It also depends on the distance between both components and the stellar wind velocity, mainly. The estimated X-ray luminosity is

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(2.6–9.1)×10^{36} \text{ erg s}^{-1} in the 3–100 keV range, assuming an isotropic emission and a distance of 6.4 kpc (Reynolds et al. 1992). Thus, the analysis of the X-ray spectrum from the neutron star at different orbital phases provides us the variability of the model parameters we can use to compare with accretion models.

The Monitor of All Sky X-ray Image (MAXI) presents both all sky coverage and moderate energy resolution which gives us the possibility to investigate the orbital light-curves and the orbital phase resolved spectra of 4U 1538–52. In this paper, we used MAXI data to study the light curves, the orbital phase averaged spectrum, and the orbital phase resolved spectra of 4U 1538–52 in the 2–20 keV energy range. Contrary to previous studies which focus on specific orbital phases at specific times, the more than five years of MAXI data, analysed in this work tend to smear out short time scale variations and give more weight to long term accretion structures. We describe the observations and the data reduction in Sect. 2, present the timing analysis in Sect. 3, the spectral analysis in Sect. 4 and in Sect. 5 we summarise our results. All uncertainties are hereafter given at the 90% ($\Delta \chi^2 = 2.71$) confidence limit, unless otherwise specified.

2. Observations

MAXI is the first astronomical mission on the International Space Station (ISS; Matsuoka et al. 2009). MAXI attached to the Japanese Experiment Module–Exposed Facility (JEM–EF). It consists in two types of X-ray slit cameras. The main X-ray camera is the Gas Slit Camera (GSC; Mihara et al. 2011) operating in the 2–20 keV energy range. The second one is the Solid-state Slit Camera (SSC; Tomida et al. 2011) operating in the 0.7–7 keV energy range. The in-orbit performance of GSC and SSC is presented in Sugizaki et al. (2011) and Tsunemi et al. (2010), respectively.

Every hour and a half MAXI covers almost the entire sky per ISS orbit. Therefore, the source is observed around 1 ks per day by MAXI. We accumulated the exposure time to extract the spectra with enough level of counts.

3. Timing analysis

We have first obtained the MAXI/GSC 1-orbit light-curve\footnote{MAXI homepage \url{http://maxi.riken.jp}} of 4U 1538–52 from MJD 55 058 to MJD 56 821 (i.e. nearly 5 years) in the 2–20 keV energy band to estimate the orbital period of the binary system. Then, we have searched for a period assuming a sinusoidal signal and the error was obtained using the relations derived by Leahy (1987). Our result was $P_{\text{orb}} = 3.7285 \pm 0.0006$ days which is consistent with the value given by Mukherjee et al. (2006) ($P_{\text{orb}} = 3.728382 \pm 0.000011$ days). Then we have folded the light curves with the best orbital period to obtain energy resolved orbital intensity profiles. The orbital phase reference is taken from Mukherjee et al. (2006), with the phase zero corresponding to mid-eclipse in Fig. 1.
Fig. 1. MAXI orbital phase resolved background subtracted light curves of 4U 1538–52 in 2–4 keV, 4–10 keV, and 10–20 keV energy ranges using orbital parameters reported by Mukherjee et al. (2006), and folded it relative to the mid-eclipse phase.

The MAXI/GSC data used in our analysis were extracted by the MAXI on-demand process using a 1.6 degrees radius circular region centred at the X-ray source position with an annulus background region where we excluded other bright sources in the field. In Fig. 1, we show the background subtracted light curves for the MAXI/GSC camera, folded at the orbital period of ~3.73 days, in the energy ranges 2–4 keV, 4–10 keV, and 10–20 keV. We note that the X-ray flux during the eclipse is compatible with no counts in all the energy bands taking the uncertainties into account.

4. Spectral analysis

4.1. Orbital phase averaged spectrum

We have extracted the orbital phase averaged spectrum of 4U 1538–52 with MAXI/GSC for the same observation duration using the MAXI on-demand processing. For spectral analysis we used the XSPEC version 12.8.1 (Arnaud (1996)) fitting package, released as a part of XANADU in the HEASoft tools. We used FTOOL GRPPHA to group the raw spectra as 2-133 channels by 4, 134-181 channels by 8, and 182-1199 channels by 12 so that all the spectral bins were Gaussian distributed. We tested both phenomenological and physical models commonly applied to accreting X-ray pulsars.

Comptonization models attempt to provide a better physical description of the underlying emission mechanisms. We fitted the 2–20 keV energy spectrum with several successful models. We have chosen Comptonization (in XSPEC terminology: Sunyaev & Titarchuk (1980)) as the representative model, i.e., a Comptonization of cool photons on hot electrons modified by an absorbing column

2 http://maxi.riken.jp/mxondem
Table 1. Fitted parameters for the MAXI spectra in Fig. 2

| Parameter          | CompaST          |
|--------------------|------------------|
| $N_H$ ($\times 10^{22}$ cm$^{-2}$) | 4.9$\pm$1.4     |
| $kT$ (keV)         | 4.7$^{+0.7}_{-0.5}$ |
| $\tau$            | 17.8$^{+2.4}_{-2.2}$ |
| Norm ($\times 10^{-2}$) | 2.7$\pm$0.6  |
| Flux (2–20 keV)    | 6.5$^{+0.3}_{-0.2}$ |
| $\chi^2$/dof      | 40.6/43          |

Notes. A CRSF was included in the model (see text). Unabsorbed flux in units of ($\times 10^{-10}$) erg cm$^{-2}$ s$^{-1}$.

along our line of sight. As MAXI/GSC does not have good enough sensitivity to detect the fluorescence iron emission line at $\sim$6.4 keV we have not needed to include it in our fits. The absorption column $N_H$ fitted with other models were consistent, within uncertainties, and we will use the CompaST values in the rest of this work. The soft excess at low energies (0.1–1 keV) presents in 4U 1538–52 and other HMXBs has not been considered here because it lies outside the energy range of the MAXI/GSC spectrum. Nevertheless, the well known fundamental cyclotron resonant scattering feature (CRSF) at $\sim$21 keV modifies the X-ray continuum above 16 keV (see references in Sect. II). Therefore, we included in our models the CRSF fixing its parameters with the values obtained by Rodes-Roca et al. (2009). We used the absorption cross sections from Verner et al. (1996), and the abundances are set to those of Wilms et al. (2000). The fitted parameters for the continuum model is listed in Table 1. In Fig. 2 we plot our data together with the absorbed CompaST best-fit model, and residuals of the fit as the difference between the observed flux and model flux divided by the uncertainty of the observed flux.

As the total Galactic H$_I$ column density in the line of sight of 4U 1538–52 is in $(9.14 \pm 9.70) \times 10^{21}$ atoms cm$^{-2}$ range (Dickey & Lockman (1990), Kalberla et al. (2005)), the neutron star is moving almost all the orbit in additional absorbing material related to the stellar wind of the companion. Therefore, we studied its variability around the orbit extracting phase resolved spectra in the next Section.

4.2. Orbital phase resolved spectra

Previous works in this sense have been carried out for only one or two orbits (Makishima et al. (1987), Robba et al. (2001), Mukherjee et al. (2006), Rodes-Roca (2007)) or for a shorter orbital phase (Rodes-Roca et al. (2011)). We have obtained orbital phase resolved spectra of the HMXB pulsar 4U 1538–52 accumulating the 60 s duration scans into six $\approx$0.5 day orbital phase bins outside of eclipse (Nakahira et al. (2012), Doroshenko et al. (2013), Islam & Paul (2014)).

In this analysis, we have fitted the orbital phase resolved spectra with the same model we used in the orbital phase averaged spectrum (see previous Section). Although we also extracted the MAXI/SSC spectra, the time coverage of the SSC was too low to study spectral and flux changes. We therefore concentrate on the GSC data. For all the spectral models, we kept the parameters free, the orbital phase resolved spectra were fitted from energy range 2 keV to 20 keV. The range of
the $\chi^2$ values was 0.7–1.1 for all fits. The data together with the absorbed CompST best-fit model, and residuals as the difference between observed flux and model flux divided by the uncertainty of the observed flux, are shown in Fig. 3. We obtained six orbital phases to keep a good S/N ratio. In Fig. 3 we plot the highest and lowest flux spectra and two sets of residuals.

The variability of the spectral parameters for the absorbed thermal Comptonization with a spherical geometry (CompST in XSPEC) is shown in Figure 4. The Comptonization parameter, $y = kT\tau^2/(m_e c^2)$, determines the efficiency of the Comptonization process (Titarchuk (1994), Prat et al. (2008)) and we derived its value from the six orbital phase ranges (see Table 2). This indicates an efficient process which corresponds to a moderate accretion rate, being higher in the orbital phase ranges just before and after the eclipse than in the mid orbital phase ranges. Moreover, it is consistent with similar observations of Vela X–1 performed by MAXI (Doroshenko et al. (2013)). The unabsorbed flux along the orbit has no significant evolution, $(7.2 - 8.8) \times 10^{39}$ erg cm$^{-2}$ s$^{-1}$ in the fitting energy range (2-20) keV, showing that the accretion rate is quite stable, in the phase bins considered.

4.3. The stellar wind in 4U 1538–52

One way to study the stellar wind in HMXBs is in terms of the variability of the equivalent hydrogen column density throughout the binary orbit. In fact, a simple spherically symmetric stellar
wind model may describe the observed orbital dependence of the column density for certain range of the orbital inclination (Clark et al. (1994), Mukherjee et al. (2006), Rodes et al. (2008)). In order to obtain an estimation of the stellar mass-loss rate, the equivalent absorption column was derived by integrating the wind density along the line of sight to the X-ray source, i.e., combining the simple spherically symmetric wind model and conservation of mass equations. Then, the hydrogen number density at distance \( r \) from the donor, \( n_H(r) \), can also be estimated through the following relationship:

\[
    n_H(r) = \frac{X_H \dot{M}}{m_H v_\infty (1 - R_*/r)^\beta / 4 \pi r^2}
\]

with \( m_H \) the hydrogen atom mass, \( X_H \) the hydrogen mass fraction, \( \dot{M} \) the mass loss rate, \( v_\infty \) the terminal velocity of the wind in the range 1400–2 800 km s\(^{-1}\) (Abbott 1982), where \( r = 23.4 \, R_\odot \) represents the binary separation (Clark 2000), the parameter \( \beta = 0.8 - 1.2 \) is the velocity gradient for an OB supergiant (Friend & Abbott 1986), and \( R_* = 17.2 \, R_\odot \) (Reynolds et al. 1992) the radius.
Fig. 4. Orbital phase changes in the free spectral parameters fitted in 2–20 keV band. Spectral parameters associated with the absorbed Comptonization of cool photons by hot electrons (\textit{compST}) model. Unabsorbed flux in units of $10^{-10}$ erg cm$^{-2}$ s$^{-1}$.

of the donor. The resultant $N_{\text{H}}$ arises from the integration of $n_{\text{H}}(r)$ along the line of sight. In Fig. 5 we show the equivalent hydrogen column density measured with MAXI as a function of the orbital phase.

The lines represent the model calculations of the absorption column density for different values of $\beta$ and $\dot{M} (M_\odot \text{yr}^{-1})/v_\infty$ (km s$^{-1}$), in units of $10^{-9}$. The red line corresponds to $\beta = 0.8$ and $\dot{M}/v_\infty = 0.3$; the blue line corresponds to $\beta = 1.2$ and $\dot{M}/v_\infty = 0.3$; the green line corresponds to $\beta = 0.8$ and $\dot{M}/v_\infty = 3.0$; and the turquoise line corresponds to $\beta = 1.2$ and $\dot{M}/v_\infty = 3.0$. Our best least-squares fit is the black line which corresponds to $\beta = 1.2$ and $\dot{M}/v_\infty = 1.3$. The value of the other orbital parameters have been taken from Clark (2000) and Mukherjee et al. (2006). As can be seen, our data can not discriminate between different values of $\beta$. We also note that our estimations of $\dot{M}/v_\infty$ did not change very much with the eccentricity and orbital inclination. None of the fits could describe all the experimental data. Consequently, we looked for the upper and lower values of a better fit. From the eclipse egress to the mid-orbit, the models with a lower $\dot{M}/v_\infty$ fit the data well (red and blue lines), but from the mid-orbit to the eclipse ingress, a higher value of $\dot{M}/v_\infty$ (green and turquoise lines) is required. This fact suggest that the neutron star is crossing an overdense region of the stellar wind and/or some trailing accreting material passing through the line of sight. Moreover,
hydrodynamical simulations show that the interaction between the gravitational field of the neutron star, the radiation field of the X-ray source and the stellar wind, may lead to the formation of large structures of higher density compared to the unperturbed stellar wind (Blondin et al. 1990). These structures are formed mainly in the trail of the neutron star. Consequently, they act as X-ray absorbers predominantly in the second half of the orbit. In Vela X–1, an excess of absorption at late orbital phases is usually observed. This excess has been successfully modelled by means of the above-mentioned hydrodynamical simulations (Manousakis et al. 2012). The fact that Vela X–1 and 4U 1538–522 are similar systems, since both have a donor of spectral type B0–0.5 I and very short orbital periods, imply that these structures are very likely to be present in 4U 1538–522.

Assuming that the eclipse egress ($\phi = [0.1 - 0.3]$) represents the unperturbed wind, the best value of $M/v_\infty$ would be in the range $(0.3 - 1.0) \times 10^{-9} \, M_\odot \, yr^{-1}/(km \, s^{-1})$. The median value of this range, namely $0.65 \times 10^{-9} \, M_\odot \, yr^{-1}/(km \, s^{-1})$, is in excellent agreement with that obtained by Clark et al. (1994). For the terminal velocities reported by Abbott (1982) this would translate into a mass loss of $(0.4 - 2.8) \times 10^{-6} \, M_\odot \, yr^{-1}$ for QV Nor.

Nieuwenhuijzen & de Jager (1990) investigated the dependence of $M$ on the stellar fundamental parameters mass $M$, radius $R$, and luminosity $L$ for a large sample of stars. They presented a simple parametrization that gave a good description of observed mass-loss rates over the whole range of parameters.
Hertzsprung-Russell diagram. We noted that their formula has different numbers in the abstract than in the text which we think it is the correct one. Therefore, we used the following expression:

$$M = 9.55 \times 10^{-15} \left( \frac{L}{L_\odot} \right)^{1.24} \left( \frac{M}{M_\odot} \right)^{0.16} \left( \frac{R}{R_\odot} \right)^{0.81} M_\odot \text{ yr}^{-1}. \quad (2)$$

The astrophysical data for QV Nor are $\log(L/L_\odot) = 5.21 \pm 0.13$, $R/R_\odot = 17.2 \pm 1.0$, and $M/M_\odot = 20 \pm 4$ (Reynolds et al. 1992). By substituting these values in the Eq. 2, the mass-loss rate was estimated as $\dot{M} = 4.5 \times 10^{-7} M_\odot \text{ yr}^{-1}$, in agreement with the smaller value obtained by considering a spherically symmetric wind.

Our values agree also with the Abbott's correlation between mass-loss rate and luminosity, $M = (0.03 - 1.0) \times 10^{-6} M_\odot \text{ yr}^{-1}$ (Abbott 1982). Finally, Clark et al. (1994), using a Ginga pointed observation, obtained a range $\dot{M} = (0.9 - 1.9) \times 10^{-6} M_\odot \text{ yr}^{-1}$, for the eclipse egress data, compatible with our estimation. Therefore, combining all the previous estimates we can deduce a mass loss range for 4U 1538–52 of $\dot{M} = (1.2 \pm 0.7) \times 10^{-6} M_\odot \text{ yr}^{-1}$. These values would point to terminal wind velocities larger than 2 000 km s$^{-1}$, closer to the upper limit (2 800 km s$^{-1}$) of the range derived by Abbott (1982).

5. Conclusions

From the analysis of the MAXI/GSC light curve, we have estimated the orbital period of the binary system, $P_{\text{orb}} = 3.7285 \pm 0.0006$ days, being in agreement with the best value derived by Mukherjee et al. (2006). From the unabsorbed flux, the X-ray luminosity in the 2–20 keV energy band was found, $L_X = (3.8 \pm 0.5) \times 10^{36}$ erg s$^{-1}$. We have investigated the long term orbital variation of spectral parameters doing an orbital phase resolved spectroscopy of 4U 1538–52 with MAXI, nearly five years of data. A higher value of $M/v_\infty$ is clearly needed to fit the absorption column from orbital phases 0.4 into eclipse ingress. Using the $N_H$ values up to orbital phase 0.3, which we consider the unperturbed wind, we measured the orbital dependence of the column density in the system and of the intrinsic source X-ray flux, and found that it was consistent on average with a spherically symmetric stellar wind from the optical counterpart of the binary system. We have estimated the mass-loss rate range of the early-type B supergiant star QV Nor to be $\dot{M} = (1.2 \pm 0.7) \times 10^{-6} M_\odot \text{ yr}^{-1}$.

The slight asymmetric accretion column distribution between eclipse-ingress and eclipse-egress seen by MAXI was also found by RXTE and BeppoSAX (Mukherjee et al. 2006), suggesting that a trailing wake of material around the neutron star seems to be a permanent structure in this source.

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