Polarization Measures and Nature of Dark Energy

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Abstract. Polarization measures, on wide angular scales, together with anisotropy data, can fix DE parameters. Here we discuss the sensitivity needed to provide significant limits. Our analysis puts in evidence that a class of models predicts low correlation or anticorrelation between polarization and anisotropy at low \( l \). This class includes open models and models with DE due to a Ratra–Peebles (RP) potential. Results on this point, given in a previous paper of ours, are updated and partially corrected. We outline that, with the sensitivity of experiments like SPOrt or WMAP, high values of \( \Lambda \) (energy scale in the RP potential) can be excluded. With the sensitivity expected for PLANCK, the selection will extend to much lower \( \Lambda \)'s.

1 Introduction

The nature of Dark Energy (DE) is one of the main puzzles of cosmology. DE was first required by SNIa data [1], but a flat Universe with \( \Omega_m \simeq 0.3 \) and \( \Omega_b h^2 \simeq 0.02 \) is also favored by CBR and LSS observations [2,3] (\( \Omega_{m,b} \): matter, baryon density parameters; \( h \): Hubble parameter in units of 100 km/s/Mpc; CBR: cosmic background radiation; LSS: large scale structure).

DE could be a false vacuum; then, its pressure and energy density (\( p_{DE} \) and \( \rho_{DE} \)) have ratio \( w = -1 \). This however requires a severe fine tuning at the end of the EW transition. Otherwise, DE can be a scalar field \( \phi \) self-interacting through a potential \( V(\phi) \) (dynamical DE [4,5,6]). Then

\[
\rho_{DE} = \frac{\dot{\phi}^2}{2a^2} + V(\phi) , \quad p_{DE} = \frac{\dot{\phi}^2}{2a^2} - V(\phi)
\] (1.1)

(derivatives are in respect to the conformal time \( t \)). As soon as \( \rho_k = \frac{\dot{\phi}^2}{2a^2} < V \), it is \( w < 0 \). For \( \rho_k/V \simeq 1/2 \), it is \( w \simeq -1/3 \) and dynamical DE approaches an open CDM behavior. Smaller \( \rho_k/V \) ratios approach \( w = -1 \) and a \( \Lambda \)CDM behavior. To work out \( w(a) \), the Friedman equations, together with the equation of \( \phi \), are to be integrated; the solutions depend on the shape of \( V \), which, in principle, is largely arbitrary.

Among potentials admitting a tracker solution, the RP [5] and SUGRA [6] expressions

\[
V(\phi) = \Lambda^{4-\alpha}/\phi^\alpha , \quad V(\phi) = (\Lambda^{4-\alpha}/\phi^\alpha) \exp(4\pi\phi^2/m_p^2)
\] (1.2)

are particularly relevant, as they originate within the frame of Supersymmetric (SUSY) theories. Here, \( \Lambda \) is an energy scale, set in the interval \( 1-10^{12} \) GeV; \( m_p \) is the Planck mass. Once \( \Lambda \) and \( \Omega_{DE} \) are fixed, the exponent \( \alpha \) is set. RP and SUGRA potentials yield fast and slowly varying \( w \), respectively.

Dynamical DE and \( \Lambda \)CDM often predict similar observational outputs. This is welcome, as \( \Lambda \)CDM is a good fit to data. However, here we show that measures of anisotropy and polarization of CBR, at large angular scales, constrain \( V \) and distinguish dynamical DE from \( \Lambda \)CDM. Accordingly, experiments in progress already exclude some parameter range for RP models and higher sensitivities can discriminate even better. In a previous paper [7], similar results were provided; that paper, however, contained a numerical mistake. Here we correct some of its quantitative results.

Our procedure includes a likelihood analysis, assuming polarization data provided by the Sky Polarization Observatory [8], both with the expected experimental noise level and a higher sensitivity. Results, however, can be straightforwardly extended to other observational contexts.

2 CMB angular spectra from the Boltzmann equations; theory

The angular CBR spectra \( C_l^{T,E,TE} \) (only \( E \)-mode is considered through this paper) can be worked out from the linear fluctuation evolution, obtained by solving the Boltzmann equation for the photon
distribution. The treatment is discussed in a number of papers [9], giving also the equations for the other components of a model. In this section we show that these equations, in several cases, yield a low or negative $C_l^{TE}$ for low $l$. All definitions used are the same as in CMBFAST [10].

This effect arises because of the simultaneous action of the ISW effect and of the opacity $\tau = \int_0^{t_o} n_e(t')\sigma_T a(t')dt'$. (Notice that $-\dot{\tau} = a n_e \sigma_T$); $n_e$ and $\sigma_T$ are the free electron density and the Thomson cross-section. ISW effect arises when we pass either from matter to curvature dominance (open models) or from matter to vacuum dominance ($\Lambda$CDM models). However, only in the former case and in the presence of opacity, anticorrelation arises. Some RP models induce anticorrelation (open models) or from matter to vacuum dominance ($\Lambda$CDM models). However, only in the former case and in the presence of opacity, anticorrelation arises. Some RP models induce anticorrelation because their features more closely approach open CDM, rather than $\Lambda$CDM.

Let us then indicate by $F_l(k,t)$ and $G_l(k,t)$ the Boltzmann components for anisotropy and polarization ($k$ is the wave–number). The equations for the $G_l$ components in flat models read:

$$
\dot{G}_l = -\dot{\tau} \left[ - G_l + \frac{\Pi}{2} (\delta l_0 + \frac{\delta l_2}{5}) \right] + \frac{k}{2l+1} [l G_{l-1} - (l+1) G_{l+1}] \quad (2.1f)
$$

($\delta_{ln}$ is the Kronecker symbol). Here

$$
\Pi = G_o + G_2 + F_2,
$$

is the only vehicle from anisotropy to polarization. In open models, eq. (2.1f) becomes:

$$
\dot{G}_l = -\dot{\tau} \left[ - G_l + \frac{\Pi}{2} (\delta l_0 + \frac{\delta l_2}{5}) \right] + \frac{\beta}{2l+1} [b_l G_{l-1} - (l+1) b_{l+1} G_{l+1}] ; \quad (2.1o)
$$

here $\beta^2 = k^2 + K$ and $b_l^2 = 1 - K l^2 / \beta^2$, with $K = -(1 - \Omega_m) H_0^2$. $H_0$ is today’s Hubble parameter.

Initially all $G_l$ terms are zero; to switch them on, the quadrupole $F_2(k,t)$ must be great when $n_e$ is not so low. For wavelengths $2\pi/k$ entering the horizon well after recombination, $F_2(k,t)$ switch on when $n_e$ has almost vanished, unless reionization occurs. Notice that the horizon size at recombination corresponds to $l \ll 200$. Without reionization, below such $l$, we expect low $G_l$’s.

Let $F_l^o$ ($G_l^o$) be the present value of harmonics. In spatially flat models, the angular spectra read

$$
C_l^T = \frac{\pi}{4} \int d^3k P_o(k) \left| F_l^o(k) \right|^2 , \quad C_l^P = \frac{\pi}{4} \int d^3k P_o(k) \left| G_l^o(k) \right|^2 , \quad C_l^{TE} = \frac{\pi}{4} \int d^3k P_o(k) F_l^o(k) G_l^o(k),
$$

$P_o(k)$ being the primeval fluctuation spectrum. In open model, instead, they read

$$
C_l^T = \frac{\pi}{4} \int d^3\beta P_o(\beta) \left| F_l^o(\beta) \right|^2 , \quad C_l^P = \frac{\pi}{4} \int d^3\beta P_o(\beta) \left| G_l^o(\beta) \right|^2 , \quad C_l^{TE} = \frac{\pi}{4} \int d^3\beta P_o(\beta) F_l^o(\beta) G_l^o(\beta) ; \quad (2.3o)
$$

here $q = (\beta^2 - 4K)^2 / (\beta^2 - K)$. Clearly, for $\Omega_m = 1$, both $\beta$ and $q$ return $k$.

Comparing eqs. (2.1) and (2.3) shows suitable shifts in the $k$–space. In particular, the $b_l$ coefficients cause a shift of $C_l$ peaks, while the passage from $k$ to $\beta$ and $q$, in eq. (2.3o), displace the power through the harmonics $F_l$ at small $l$, i.e. on scales comparable with the curvature scale. Apart of these shifts, the gravitational field fluctuations ($\dot{h}$ and $\eta$) obey similar but different equations:

**Open models**:

$$
2k^2 \dot{\eta} = 8\pi G a^2 \left[ \sum_c (\rho_c + p_c) \theta_c - \dot{\theta}_o,cr (1 - \Omega_m) / a^2 \right]. \quad (2.4o)
$$

**Flat models**:

$$
2k^2 \dot{\eta} = 8\pi G a^2 \left[ \sum_c (\rho_c + p_c) \theta_c + (\rho_{DE} + p_{DE}) \theta_{DE} \right]. \quad (2.4f)
$$

Here $\sum_c$ sums over all (relativistic or non-relativistic) matter components apart of DE; $\rho_{o,cr}$ is today’s critical density. The wave–numbers $\tilde{k}^2$ and $k^2$ are just shifted by $3K$.

Notice that, in a $\Lambda$CDM model, no DE fluctuations exist while $\rho_{DE} = -p_{DE}$, so that the second term at the r.h.s. of eq. (2.4f) vanishes. In models with dynamical DE, instead, $\theta_{DE} \neq 0$ and $\rho_{DE} \neq -p_{DE}$. Accordingly, the second term in square brackets in eq. (2.4f) may read $\theta_{DE} \rho_{o,cr} (1 - \Omega_m) (1 + w) (\rho_{DE}/\rho_{o,DE})$ (the last parenthesis tells us how DE energy scales with $a$). This term is analogous to the second term in square bracket in eq. (2.4o) and would coincide with it if $w = -1/3$ and, namely, if $\theta = -\dot{h}/2$. If this is true, apart of a different power distribution along the $l$ axis and
Figure 1: TE correlation spectra $C_l = l(l+1)C_{l}^{TE}/4\pi$ for ΛCDM (the same in the three plots) and RP. Solid (dashed) lines refer to $\tau = 0.14$ ($0.20$). In the third plot, low–l peaks are negative.

geometric effects at greater $l$, there can be similarities in the behavior of open and dynamical DE models. The relation between $\dot{h}/2$ and $\theta_{DE}$ can then be studied through the equation

$$\theta_{DE} + \frac{\dot{h}}{2} = -\frac{1}{1+w}[\delta_{DE} + \frac{3}{a}(c_s^2 - w)\delta_{DE}],$$

(2.5)

whose validity indicates that DE behaves as a fluid. An order of magnitude estimate, however, tells us soon that $\theta_{DE} \sim k^2 t \delta_{DE}$; then, the ratio between $\theta_{DE}$ and the r.h.s. is $\sim (t/L)^2$, where $L$ is the scale related to $k$. Before horizon crossing ($t \ll L$), the $\theta_{DE}$ term is negligible, in comparison with the r.h.s. Hence, $\dot{h}/2$ equates the r.h.s. and, therefore, the ratio $-\dot{h}/2\theta_{DE}$ exceeds unity. At horizon crossing such ratio must approach unity and keep such as $t$ grows greater than $L$. The main differences between open and dynamical DE models, in the r.h.s. of eqs. (2.4) are therefore relegated to times before horizon crossing. Afterward, the residual difference is due to a factor $1 + w$.

The changes in $\eta$ directly act on $F_2$. The equation fulfilled by this spectral component reads:

$$\dot{F}_2(k,t) = -\dot{\tau} \left[-F_2(k,t) + \frac{\Pi(k,t)}{10}\right] + \frac{k}{5}[2F_1(k,t) - 3F_3(k,t)] + \frac{8}{5} \left[\dot{\eta}(k,t) + \frac{\dot{h}(k,t)}{6}\right],$$

(2.6f)

in flat models, while in open models we have:

$$\dot{F}_2(\beta,t) = -\dot{\tau} \left[-F_2(\beta,t) + \frac{\Pi(\beta,t)}{10}\right] + \frac{\beta}{5}[2b_2 F_1(\beta,t) - 3b_3 F_3(\beta,t)] + \frac{8}{5} \frac{\beta}{k} b_2 \left[\dot{\eta}(\beta,t) + \frac{\dot{h}(\beta,t)}{6}\right].$$

(2.6o)

Accordingly, when open models show negative TE correlations at low $l$, we expect something similar in flat RP models. On the contrary, models like SUGRA, with a cosmic acceleration closer to ΛCDM models, are not expected to give negative TE correlation.

Let us then discuss how the likelihood distribution can be obtained from model spectra, taking into account that the number of pixels for anisotropy and polarization ($N_T$ and $N_P$) in our (artificial) data are different. Let $T_j$ be the anisotropy data measured in $N_T$ pixels and $Q_j$ and $U_j$ be the Stokes parameters in $N_P$ pixels. In general, let be $x \equiv (T_1, ..., T_{N_T}, Q_1, ..., Q_{N_P}, U_1, ..., U_{N_P})$. $x$ is a vector of $N_s = N_T + 2N_P$ components, defining an observed state of anisotropy and polarization. Once a model is assigned, the $C_l$ are uniquely determined. Passing to a data vector $x$, instead, amounts to performing a model realization. Vice–versa, once the $N_s$ component data vector $d$ is given, the model is not uniquely fixed.
Then, the likelihood of a model, whose angular spectra are $C_l$, when the data $d$ are observed, reads

$$L(d|C_l^A) \propto [\det M]^{-\frac{1}{2}} \exp\left[-\frac{1}{2} d^T M^{-1} d\right].$$

The main ingredient of $L$ is the correlation matrix $M_{ij} = \langle x_i^T x_j \rangle = S_{ij} + N_{ij}$; here $S_{ij}$ is the signal term and $N_{ij}$ is due to the noise. The components $M_{ij}$ yield the expected correlation between the $i$th and $j$th elements of data vectors $x$ corresponding to particular choices of $C_l$ [10]. The construction of the noise term is simpler, as we expect no noise correlation, and the matrix $N_{ij} = \delta_{ij} \sigma_{T,pix}^2$ (for $i = 1, ..., N_T$) and $N_{ij} = \delta_{ij} \sigma_{P,pix}^2$ (for $i = N_T + 1, ..., N_s$) is diagonal.

### 3 Angular spectra and likelihood: results.

In Fig. 1 we show the spectra $C_l^{TE}$ for ΛCDM and RP models, if $\tau = 0.14$ and $\tau = 0.20$ (WMAP[11] suggests that $\tau = 0.17 \pm 0.04$). The value of $\lambda = \log_{10}(\Lambda/\text{GeV})$ is given in top of the frames. Low-$l$ anticorrelation is found only for $\lambda > \sim 10$ and is shown in the third plot. The differences between ΛCDM and RP increase for growing $\lambda$, but are already significant even for $\lambda = 2$.

At low $l$, cosmic variance must be taken into account, aside of instrumental variance. We then perform a large number (1000) of realizations of RP sky models for $\lambda = 2, 5$ and 8, for $\tau = 0.14$ and 0.20. In Fig. 2 (3) we report results for $\tau = 0.14$ (0.20). The values of $\lambda$ are shown in top of each plot. Lower and upper plots correspond to $\sigma = 2$ and 0.2 $\mu$K, respectively. The former value approaches WMAP and SPOrt expectations; the latter value might be approached by the PLANCK [12] experiment.

### 4 Conclusions

These figures show that, as expected, RP models are more easily distinguishable from ΛCDM for lower $\sigma$ and higher $\lambda$. We can assume that a RP model gives a signal different from ΛCDM when $\lambda > 0$ is detected. However, even in the less favorable case considered, when $\sigma = 2\mu$K, $\tau = 0.14$ and $\lambda = 2$, the peak likelihood is at $\lambda > 0$ in $\sim 72$% of cases. For $\lambda = 5$, this fraction reaches $\sim 91$%. Likelihood distributions, averaged both over cosmic and instrumental variances, tell us that RP models begin to
be distinguishable from $\Lambda$CDM, at $1-\sigma$ level, for $\lambda \sim 5$ and $\tau = 0.20$, if $\sigma = 2\mu K$. For $\sigma = 0.2\mu K$, instead, this is already true for $\lambda = 2$. See the figures for further details.

A general conclusion that can be drawn from this analysis is that cosmic variance is far from being a serious limit to model parameter detection from large angle spectral analysis, at the present sensitivity levels. Even for a sensitivity improved by a factor 10, histograms are still contained inside the curve, showing that there is a relevant space for further improvements.

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