Analytical relations concerning the collapse time in hierarchically clustered cosmological models

M. Gambera\textsuperscript{1,2}, A. Pagliaro\textsuperscript{1}

\textsuperscript{1} Istituto di Astronomia dell’Università di Catania, Viale A.Doria, 6 - I 95125 Catania, ITALY

\textsuperscript{2} Osservatorio Astrofisico di Catania and CNR-GNA, Viale A.Doria, 6 - I 95125 Catania, ITALY

Abstract. By means of numerical methods, we solve the equations of motion for the collapse of a shell of baryonic matter, made of galaxies and substructure falling into the central regions of a cluster of galaxies, taking into account the effect of the dynamical friction. The parameters on which the dynamical friction mainly depends are: the peaks’ height, the number of peaks inside a protocluster multiplied by the correlation function evaluated at the origin, the filtering radius and the nucleus radius of the protocluster of galaxies. We show how the collapse time $\tau$ of the shell depends on these parameters. We give a formula that links the dynamical friction coefficient $\eta$ to the parameters mentioned above and an analytic relation between the collapse time and $\eta$. Finally, we obtain an analytical relation between $\tau$ and the mean overdensity $\bar{\delta}$ within the shell. All the analytical relations that we find are in excellent agreement with the numerical integration.

Key words: cosmology: theory-large scale structure of Universe - galaxies: formation

Send offprint requests to: M. Gambera E-mail mga@sunct.ct.astro.it
1. Introduction

The problem of the formation and evolution of clusters of galaxies has been one of the crucial topics of the last years (see e.g. Ryden & Gunn 1987, Colafrancesco et al. 1989, Antonuccio-Delogu 1992, Kaiser 1993, Colafrancesco & Vittorio 1993, Croft & Efstathiou 1994, Dutta & Spergel 1994, Mosconi et al. 1994, Sutherland & Dalton 1994, Efstathiou 1994 and Colafrancesco et al. 1995). It is well known that the formation of cosmic structures is strictly related to the evolution of the density perturbations: in the present paradigm of structure formation, it is generally assumed that cosmic structures of size $\sim R$ form preferentially around the local maxima of the primordial density field, once it is smoothed on the filtering scale $R_f$. These linear density fluctuations eventually evolve towards the nonlinear regime under the action of gravitational instability; they detach from the Hubble flow at turn around epoch $t_m$, given by:

$$t_m = \left[ \frac{3\pi}{32G\rho_b}(1 + \bar{\delta}) \right]^{1/2} (1 + z)^{3/2}$$

where $\rho_b$ is the mean background density, $z$ is the redshift and $\bar{\delta}$ is the mean overdensity within the nonlinear region. After the turn around epoch, the fluctuations start to recollapse when their overdensity defined by

$$\delta(x) = \frac{\rho(x) - \rho_b}{\rho_b}$$

reaches the value $\bar{\delta} = 1$.

The evolution of the density fluctuations is described by the power spectrum, given by:

$$\mathcal{P}(k) \equiv \langle |\delta_k|^2 \rangle$$

with:

$$\delta_k \equiv \int d^3k e^{-i\mathbf{k}\cdot\mathbf{x}} \delta(x)$$

Since the density field depends on the power spectrum, which in turn depends on the matter that dominates the universe, the mean characteristics of the cosmic structures depend on the assumed model. In this context the most successful model is the biased Cold Dark Matter (hereafter CDM) (see e.g. Kolb & Turner 1990; Peebles 1993; Liddle & Lyth 1993) based on a scale invariant spectrum of density fluctuations growing under gravitational instability. In such scenario the formation of the structures occurs through a "bottom up" mechanism. A simple model that describes the collapse of a density perturbation is that by Gunn & Gott (1972, hereafter GG72). The main assumptions of this model are: (a) the symmetry of the collapse is spherical; (b) the matter distribution is
uniform in that region of space where the density exceeds the background density; (c) no tidal interaction exists with the external density perturbations and (d) there is no substructure (collapsed objects having sizes less than that of main perturbation).

Point (d) is in contradiction to the predictions of CDM models. It is well known that in a CDM Universe, an abundant production of substructures during the evolution of the fluctuations is predicted.

The problem of the substructures in a CDM Universe and his consequences on structure formation have been widely studied in previous papers (see e.g. Antonuccio-Delogu 1992, hereafter A92; Antonuccio-Delogu & Atrio-Barandela 1992, hereafter AA92; Antonuccio-Delogu & Colafrancesco 1994, hereafter AC94, Del Popolo et al. 1996, Gambera 1997 and Del Popolo & Gambera 1997 hereafter DG97). Being a rather recent topic in cosmology much work has still to be done.

The presence of substructure is very important for the dynamics of collapsing shells of baryonic matter made of galaxies and substructure of $10^6 M_\odot \div 10^9 M_\odot$, falling into the central regions of a cluster of galaxies. As shown by Chandrasekhar & von Neumann (1942, hereafter CvN42; 1943), in the presence of substructure it is necessary to modify the equation of motion:

$$\frac{d^2 r}{dt^2} = - \frac{GM}{r^2(t)}$$

since the graininess of mass distribution in the system induces dynamical friction that introduces a frictional force term.

Adopting the notation of GG72 (see also their Eqs. 6 and 8) and remembering that $T_{co}/2$ is the collapse time of a shell of baryonic matter in the absence of dynamical friction (GG72), one can write:

$$T_{co} = \frac{\pi \bar{\rho}_i \rho_{ci}^{1/2}}{H_i(\bar{\rho}_i - \rho_{ci})^{3/2}}$$

where $\rho_{ci}$ is the critical density at a time $t_i$ and $\bar{\rho}_i$ is the average density inside $r_i$ at $t_i$. The equation of motion of a shell of baryonic matter in presence of dynamical friction (Kandrup 1980, hereafter K80; Kashlinsky 1986 and AC94), using the dimensionless time variable $\tau = t/T_{co}$, can be written in the form:

$$\frac{d^2 a}{d\tau^2} = - \frac{4\pi G \rho_{ci}(1 + \bar{\delta}_i)}{a^2(t)} T_{co} - \eta T_{co} \frac{da}{d\tau}$$

where $\bar{\delta}_i$ is the overdensity within $r_i$, $\eta$ is the coefficient of dynamical friction and $a(r_i, t)$ is the expansion parameter of the shell (see GG72 Eq. 6), that can be written as:

$$a(r_i, t) = \frac{r(r_i, t)}{r_i}$$
Supposing that there is no correlation among random force and their derivatives, we have:

\[
\eta = \int \frac{d^3 F W(F) F^2 T(F)}{2 \langle v^2 \rangle}
\]  

(K80), where \( T(F) \) is the average "duration" of a random force impulse of magnitude \( F \), \( W(F) \) is the probability distribution of stochastic force (which for a clustered system is given in Eq. 37 of AA92).

DG97 solved Eq. (7) numerically, showing qualitatively how the expansion parameter \( a(\tau) \) depends on the dynamical friction coefficient and how \( \tau \) changes in the presence of dynamical friction, but without undertaking a more complete study of the dependence on the parameters.

The plan of this paper is as follows. In Sect. 2 we show how \( \tau \) depends on the peaks' height \( \nu_c \), on the parameter \( \Xi \) that we define there as the correlation function at the origin multiplied by the total number of peaks inside a protocluster, on the filtering radius \( R_f \) and on the nucleus radius of the protocluster \( r_0 \). A more detailed description of these parameters is given below. In Sect. 3 we give an analytical relation between the dynamical friction coefficient and the collapse time:

\[
\tau = f_1(\eta)
\]  

(10)

In Sect. 4 we give a semi-analytical relation that links the dynamical friction coefficient with the parameters on which it depends:

\[
\eta = f_2(\nu_c, R_f, r_0, \Xi)
\]  

(11)

Linking these two relations we also find the dependence of the dimensionless collapse time \( \tau \) on the parameters used:

\[
\tau = f_1 \circ f_2 \equiv f_1[f_2(\nu_c, R_f, r_0, \Xi)]
\]  

(12)

In Sect. 5 we give a semi-analytical relation between \( \tau \) and \( \delta \) and \( \eta \):

\[
\tau = f_3(\delta, \eta)
\]  

(13)

Finally, in Sect. 6 we summarize our results and comment on their possible implications.

2. The collapse time

DG97 showed how the expansion parameter \( a(\tau) \) depends on the dynamical friction, solving Eq. (7) by means of a numerical method but not taking into account the parameters on which \( \eta \) depends.
Here we examine how the dynamical friction coefficient $\eta$ varies according to the parameters and how the collapse time depends on them. We consider Eq. (9) and the functions $T(F)$ and $W(F)$, the former given by Chandrasekhar (1943), the latter (for clustered system) given by the so-called Holtsmark law (Holtsmark 1919):

$$W(F) = \frac{2F}{\pi} \int_0^{\infty} dkk \sin(kF) \exp \left[ -\frac{3}{2} N \left( \frac{Gm_{typ} k}{r_{typ}} \right)^{3/2} \right]$$

$$= 4\pi |F| W(|F|)$$  \hspace{1cm} (14)

where $W(F)dF$ is the probability for a test particle of experiencing a force in the range $F \leq F + dF$, $N$ is the total number of particles, $m_{typ}$ is a typical particle mass and $r_{typ}$ is a typical distance among the particles.

According to K80:

$$W(|F|) = \frac{1}{2\pi^2 |F|} \int_0^{\infty} dkk \sin(k|F|) A_f(k)$$  \hspace{1cm} (15)

with:

$$A_f(k) = \lim_{N \to \infty} A N(k) = A_{nor} \exp(\mathcal{F})$$  \hspace{1cm} (16)

The original expression for $\mathcal{F}$ given by K80 has been modified by AA92 to take into account clustering, and turns out to be given by:

$$\mathcal{F} = \left\{ A_{uncl}(k) + A_{cl}(k) \left[ 1 + \frac{\Sigma_{cl}(k)}{2A'_{\Sigma,cl}} \right] \right\}$$  \hspace{1cm} (17)

where $A_{nor}$, $A_{cl}(k)$, $A_{uncl}(k)$, $A'_{\Sigma,cl}(k)$ and $\Sigma_{cl}(k)$ are given in AA92 respectively by the Eqs. (21), (29), (31), (32) and (36).

Here, we want to remind that $\Sigma_{cl}(k)$ is a linear function of the correlation function $\xi(r)$ and that the general expression for $\Sigma_{cl}(k)$, adopting the notation of AA92 is given by:

$$\Sigma_{cl}(k) = \int_{\mathbb{R}^3} d^3r_1 \int_{\mathbb{R}^3} d^3r_2 \exp \left( i \frac{Gm_{av} k \vec{r}_1}{r_1^2} \right) \cdot \exp \left( i \frac{Gm_{av} k \vec{r}_2}{r_2^2} \right) \tau_{cl}(r_1) \tau_{cl}(r_2) \xi(|r_1 - r_2|)$$  \hspace{1cm} (18)

where masses are measured in units of solar mass $M_{\odot}$ and distances in $h^{-1}$ Mpc, so that $k$ will be measured in unit of $h^{-1}(\text{Mpc})^2/GM_{\odot}$, with $m_{av}$ the average mass of the substructure. Since we have $m_{av} \geq 10^6 M_{\odot} >> 1 M_{\odot}$ for all the cases that we consider in...
this paper, we can adopt the asymptotic expansion of $\Sigma_{cl}$ demonstrated in the appendix of AA92:

$$\Sigma_{cl}(k) \propto \xi(0)$$  \hspace{1cm} (19)

(see AA92 for details). $\Sigma_{cl}$ does not depend on $\xi(r)$, but only on $\xi(0)$ (that is $\xi(r)$ calculated at the origin).

We solve Eq. (9) and the other equations related for a outskirts shell of baryonic matter with $\bar{\delta} = 0.01$ inside the spherical regions (protocluster), for different values of $\nu_c, R_f, r_0$ and $\Xi$ (where $\Xi \equiv N_{tot} \cdot \xi(0)$), the latter quantity being better defined in Sect. 4.

After having determined $\eta$ solving numerically Eq. (9), we get $\tau$ as a function of $\nu_c, R_f, r_0$ and $\Xi$ solving Eq. (7). We perform these calculations for different set of values of $\nu_c, R_f, r_0$ and $\Xi$ inside the following intervals:

\begin{align*}
1.2 & \leq \nu_c \leq 3.2 \\
0.5 \, h^{-1}\text{Mpc} & \leq r_0 \leq 10 \, h^{-1}\text{Mpc} \\
10^{-3} \, h^{-1}\text{Mpc} & \leq R_f \leq 1 \, h^{-1}\text{Mpc} \\
20 & \leq \Xi \leq 2 \cdot 10^3
\end{align*}  \hspace{1cm} (20)

The results that we have obtained are shown in Figs. 1-4. Before commenting upon the figures, we want to remark that the dependence of $\tau$ on $\bar{\delta}$ is qualitatively shown in Fig. 5 by AC94. We observe that for $\bar{\delta} > 10^{-2}$ the collapse time in the presence of a dynamical friction is always larger than in the unperturbed case but the magnitude of the deviation is negligible for larger $\bar{\delta}$, whilst for $\bar{\delta} \leq 10^{-2}$ the deviations increase steeply with lower $\bar{\delta}$. Then, having considered $\bar{\delta} = 0.01$, the estimation we get for $\tau$ in Sect. 4 must be considered as a lower limit.

In Fig. 1 we show the collapse time in the presence of dynamical friction, versus the peaks' height, for different values of $\Xi$. In this picture, we show how $\tau$ grows for larger values of $\nu_c$ and for larger values of $\Xi$. Similarly, in Fig. 2 we note how $\tau$ increases for larger values of $\nu_c$ and of $R_f$. The slope of the curves confirm our prevision on the behaviour of the collapse of a shell of baryonic matter falling into the central regions of a cluster of galaxies in the presence of dynamical friction: the dynamical friction slows down the collapse (as DG97 had already shown) and the effect, as we are showing in Figs. 1 and 2, increases as $\Xi, R_f, \nu_c$ grow.
Fig. 1. Collapse time $\tau$ of a shell of matter made of galaxies and substructure when dynamical friction is taken into account, versus $\nu_c$. We assume a nucleus radius of $r_0 = 1h^{-1}\text{Mpc}$ and a filtering radius $R_f = 0.74h^{-1}\text{Mpc}$. Open circles: $\Xi = 10^3$; filled circles: $\Xi = 10^2$.

Here we want to remind that we are considering only the peaks of the local density field with central height $\nu$ larger than a critical threshold $\nu_c$. This latter quantity is chosen to be the threshold at which $r_{\text{peak}} (\nu \geq \nu_c) << l_{\text{av}}$ where $r_{\text{peak}}$ is the typical size of the peaks and $l_{\text{av}}$ is the average peak separation (see also Bardeen et al. 1986).

In Figs. 3 and 4 we show how the collapse time varies with the nucleus radius of the protocluster $r_0$. Note how $\tau$ grows as $r_0$ decreases: the smaller the nucleus of the protocluster, the larger the time of collapse in the presence of dynamical friction; besides we show how this effect increases for larger values of both $\nu_c$ and $\Xi$.

3. How the collapse time depends on the dynamical friction

Like DG97, we have solved Eq. (1) numerically and calculated the collapse time $\tau$ for different values of the dynamical friction coefficient $\eta$. However, DG97 showed the dependence of $\tau$ on $\eta$ qualitatively, whilst we have found also an analytic relation $\tau = f_1(\eta)$ that links $\tau$ with $\eta$ in the range of values $0 \leq \eta \leq 3.1$ (see also Gambera 1997). The
results of our calculations are shown in Fig. 5, where we report the collapse time \( \tau \) versus the dynamical friction coefficient. We have determined the function \( \tau = f_1(\eta) \) through a least squares fit method. Our best solution is given by:

\[
\tau = a_1 \eta^2 + a_2 \eta + a_3 \tag{21}
\]

with:

\[
a_1 = 0.00595923,
\]

\[
a_2 = 0.0493460,
\]

\[
a_3 = 0.492164,
\]

Eq. (21) and our results for \( 0 < \eta < 3.1 \) from the numerical integration of Eq. (7) are in excellent agreement. As a matter of fact, the correlation coefficient value obtained is:

\[
r^2 = 1.000.
\]
Fig. 3. Collapse time $\tau$ versus $r_0$. We assume a filtering radius $R_f = 0.74 h^{-1}$ Mpc. and a total number of peaks of substructure $\Xi = 10^3$. Open circles: $\nu_c = 3$; crosses: $\nu_c = 2$; filled circles: $\nu_c = 1.24$.

4. How the collapse time depends on the parameters

As already shown by AC94, it is possible to write down the dynamical friction coefficient $\eta$ as the sum of two terms:

$$\eta = \eta_0 + \eta_{cl}$$ (22)

where $\eta_0$ is the coefficient of dynamical friction of an unclustered distribution of field particles whilst $\eta_{cl}$ takes into account the effect of clustering. We rewrite Eq. (22) as:

$$\eta = \eta_0 + \eta_{cl} = \eta_0 \left(1 + \frac{\eta_{cl}}{\eta_0}\right)$$ (23)

where, as demonstrated by A92, the ratio $\frac{\eta_{cl}}{\eta_0}$ depends only on $\Xi$ and $r_0$ whilst $\eta_0$ is given by A92:

$$\eta_0 = \frac{3.33}{\pi} \frac{[G(m)_{av} p(\gamma, \nu_c)]^{1/2}}{\langle n \rangle_{av} \langle R_{cl}^2 \rangle^{1/2}} \log(1.36 \langle n \rangle_{av}^{2/3} R_{cl}^2 \pi^{2/3}) \frac{1}{a^{3/2}}$$ (24)
Fig. 4. Collapse time \( \tau \) of a shell of matter made of galaxies and substructure when dynamical friction is taken into account, versus \( r_0 \). We assume a filtering radius \( R_f = 0.74h^{-1} \) Mpc and a peaks’ height \( \nu_c = 3 \). Open circles: \( \Xi = 10^3 \); filled circles: \( \Xi = 10^2 \).

and, for a fixed value of \( \delta \), depends only on \( \nu_c \) and \( R_f \). In the previous formula, \( c_p \) is defined in Bower (1991), whilst \( p(\gamma, \nu_c), \langle m \rangle_{av} \) and \( \langle n \rangle_{av} \) are defined in AC94. Therefore, we rewrite Eq. (23) as:

\[
\eta = \eta_0(\nu_c, R_f) \left( 1 + \frac{\eta_{cl}}{\eta_0}(r_0, \Xi) \right)
\]

(25)

where the parameters on which \( \eta \) depends are the following:
- \( \nu_c \) is the peaks’ height;
- \( R_f \) is the filtering radius;
- \( r_0 \) is the parameter of the power-law density profile. A theoretical work (Ryden 1988) suggests that the density profile inside a protogalactic dark matter halo, before relaxation and baryonic infall, can be approximated by a power-law:

\[
\rho(r) = \frac{\rho_0 r_0^p}{r^p}
\]

(26)

where \( p \approx 1.6 \) on a protogalactic scale.
Fig. 5. Collapse time $\tau$ of a shell of matter made of galaxies and substructure when dynamical friction is taken into account, versus $\eta$.

$\Xi$ is the product $N_{tot} \cdot \xi(0)$ where $N_{tot}$ is the total number of peaks inside a protocluster and $\xi(0)$ is the correlation function calculated in $r = 0$. AA92 have demonstrated that in the hypothesis $m_{av} \gg 1M_\odot$, where $m_{av}$ is the average mass of the subpeaks, the dependence of the dynamical friction coefficient on $N_{tot}$ and $\xi(r)$ may be expressed as a dependence on a single parameter that we define as:

$$
\Xi \equiv N_{tot} \cdot \xi(0) \quad (27)
$$

The analytic relation $\eta = f_2(\nu_c, R_f, r_0, \Xi)$, that links the dynamical friction coefficient with the parameters on which it depends, can be rewritten as the product of two functions:

$$
\eta = f_2(\nu_c, R_f, r_0, \Xi) = f'_2(\nu_c, R_f) \cdot f''_2(r_0, \Xi) \quad (28)
$$

where

$$
\eta_0 = f'_2(\nu_c, R_f) \quad (29)
$$

and
\[ 1 + \frac{\eta_{cl}}{\eta_0} = f''_2(r_0, \Xi) \]  

(30)

With a least square method, we find the best function \( \eta = f_2(\nu_c, R_f, r_0, \Xi) \). First, we find an analytical relation between the dynamical friction coefficient in the absence of clustering \( \eta_0 \) and the parameters \( R_f \) and \( \nu \). We obtain:

\[
\eta_0 = b_1 + b_2 \nu + b_3 R_f + b_4 \nu R_f + b_5 \nu^2 \\
+ b_6 R_f^2 + b_7 \nu R_f^2 + b_8 \nu^2 R_f^2 + b_9 \nu^2 R_f^2
\]  

(31)

with:

\[
b_1 = 1.00655066, \quad b_2 = -0.01224778, \\
b_3 = 0.03836632, \quad b_4 = -0.03337465, \\
b_5 = 0.00288929, \quad b_6 = -0.02622193, \\
b_7 = 0.04374888, \quad b_8 = 0.00690435, \\
b_9 = 0.00860054,
\]

We perform a \( \chi^2 \) test between the values of \( \eta_0 \) obtained from the last equation and the values of \( \eta_0 \) calculated integrating Eq. (24). Our result for the range \( 1.2 \leq \nu_c \leq 3.2 \) is excellent: \( \chi^2 \approx 10^{-8} \).

We do the same job for the quantity \( f''_2(r_0, \Xi) \), finding:

\[
f''_2 = c_1 + c_2 r_0 + c_3 \Xi + c_4 \Xi r_0 + c_5 r_0^2 + c_6 \Xi r_0^2
\]  

(32)

with:

\[
c_1 = 2.26516888, \quad c_2 = -0.09679056, \\
c_3 = 0.00054051, \quad c_4 = -0.0001255044, \\
c_5 = 0.00850948, \quad c_6 = 0.000009336723
\]

An analogue \( \chi^2 \) test gives \( \chi^2 \approx 10^{-8} \) for the range \( 20 \leq \Xi \leq 2 \cdot 10^3 \).

The function \( \eta = f_2(\nu, R_f, r_0, \Xi) \) is given by the product of Eqs. (31) and (32). These contain 54 terms.

\[
\eta = f_2(\nu, R_f, r_0, \Xi) = \\
= (b_1 + b_2 \nu + b_3 R_f + b_4 \nu R_f + b_5 \nu^2 + b_6 R_f^2 + b_7 \nu R_f^2 + b_8 \nu^2 R_f^2 + b_9 \nu^2 R_f^2) \\
+ b_6 \nu^2 R_f^2 + b_9 \nu^2 R_f^2) \cdot (c_1 + c_2 r_0 + c_3 \Xi \\
+ c_4 \Xi r_0 + c_5 r_0^2 + c_6 \Xi r_0^2)
\]  

(33)
However, we have also found an empirical formula with only 13 terms that represents a good approximation. We performed a $\chi^2$ test between the results obtained from the 13-term equation and the results obtained from the numerical integration. The result is $\chi^2 = 2.17 \cdot 10^{-3}$. The same test performed on the 54-term equation gives $\chi^2 \approx 10^{-8}$. Our 13-term equation reads as:

$$\eta \approx d_1 + d_2 \nu + d_3 R_f + d_4 r_0 + d_5 \Xi + d_6 \nu R_f + d_7 \nu^2 + d_8 R_f^2 + d_9 r_0^2 + d_{10} r_0 \Xi + d_{11} \nu R_f^2 + d_{12} \nu^2 R_f + d_{13} r_0^2 \Xi$$

with:

$d_1 = 2.280000$, $d_2 = -0.027743$,
$d_3 = 0.086906$, $d_4 = -0.097425$,
$d_5 = 5.440520$, $d_6 = -0.075599$,
$d_7 = 0.006545$, $d_8 = -0.059397$,
$d_9 = 0.008565$, $d_{10} = 1.26 \cdot 10^{-4}$,
$d_{11} = 0.099099$, $d_{12} = 0.015640$,
$d_{13} = 9.39 \cdot 10^{-6}$

The function $\tau = f_1 \circ f_2$ is given by:

$$f_1 \circ f_2 = a_1 [f_2(\nu, R_f, r_0, \Xi)]^2 + a_2 [f_2(\nu, R_f, r_0, \Xi)] + a_3$$

that is:

$$\tau = a_1 \left[ \eta_0 \left(1 + \frac{\eta_1}{\eta_0} \right) \right]^2 + a_2 \left[ \eta_0 \left(1 + \frac{\eta_1}{\eta_0} \right) \right] + a_3$$

where the values of $a_n$ are given in Sect. 2 and the function $f_2$ is given by Eq. (33) or by Eq. (34).

5. A semi-analytical relation between $\tau$ and $\overline{\delta}$

Our aim is to find a semi-analytical relation $\tau = f_3(\overline{\delta}, \eta)$ for the intervals $10^{-4} \leq \overline{\delta} \leq 10^{-2}$ and $0 \leq \eta \leq 3.1$.

The first step is the determination of a function $\tau = g(\overline{\delta})$ for a fixed value of $\eta$. We solve Eq. (8) for $10^{-4} \leq \overline{\delta} \leq 10^{-2}$ and $\eta = 0.01$ and by a least square method we find the function $\tau = g(\overline{\delta})$:

$$\tau = e_1 \overline{\delta}^2 + e_2 \overline{\delta} + e_3$$
with:

\[ e_1 = 5678.65, \]
\[ e_2 = 45.3429, \]
\[ e_3 = -0.01591, \]

The value of the correlation coefficient between this function and the numerical integration is \( r^2 = 1.000 \). With this method we find \( \tau = g(\delta) \) for several values of \( \eta \) inside the interval \( 0 \leq \eta \leq 3.1 \). We can write the dimensionless collapse time as the product:

\[ \tau = g(\delta) \cdot h(\eta) \]

(38)

where the function \( h(\eta) \) can be written as

\[ h(\eta) = K \cdot [1 + f(\eta)] \]

(39)

and, for \( 10^{-4} \leq \delta \leq 10^{-2} \) and \( 0 \leq \eta \leq 3.1 \), \( K \) is constant: \( K = -1494.7705898 \). So:

\[ \tau = F(\delta, \eta) = K \cdot [1 + f(\eta)] \cdot g(\delta) \]

(40)

A \( \chi^2 \) test performed between the values obtained for \( \tau \) from the numerical integration and the values obtained from Eq. (7) gives the result: \( \chi^2 \sim 10^{-4} \).

6. Conclusions and discussion

In Sect. 2 of this work we showed in a quantitative way how the collapse time \( \tau \) of a shell of baryonic matter made of galaxies and substructure depends on some parameters. When one of the parameters \( \Sigma \) or \( R_f \) or \( \nu_c \) increases, the collapse time grows. It means that the effects of the presence of dynamical friction should be more evident in the outer regions of rich clusters of galaxies. Besides, we show how the collapse time of an infalling shell increases with decreasing values of \( r_0 \), and becomes very large for \( r_0 \leq 2h^{-1} \) Mpc (see Fig. 4). As a consequence, the slowing down of the collapse of an outer shell within a cluster of galaxies owing to the dynamical friction is more remarkable in the clusters with nucleus of little dimension.

In Sect. 3 of this paper we give an analytic relation that links the dimensionless collapse time \( \tau \) with the coefficient of dynamical friction \( \eta \). This relation is in excellent agreement with the numerical integration of Eq. (7) for \( 0 \leq \eta \leq 3.1 \) (see Fig. 5). Then, we find an analytic relation between \( \eta \) and the parameters on which it depends for the value \( \delta = 10^{-2} \) (we remind that the effects of the dynamical friction are negligible for \( \delta > 10^{-2} \)). This is Eq. (28), that can be considered as a "low order" approximation to a more realistic
situation of an outer shell of a cluster of galaxies with $\delta \leq 10^{-2}$. We also find an empirical formula, Eq. (34), that is a good approximation of Eq. (33). Moreover, the dependence of the dimensionless collapse time on $\nu_c$, $r_0$, $\Xi$ and $R_f$ is shown (Eq. (35)). Finally, we give an analytical relation that links $\tau$ with $\eta$ and $\delta$. Here, we wish to stress the usefulness of an analytic relation like $\tau = K \cdot [1 + (f_1 \circ f_2)] \cdot g(\delta)$. This is a powerful tool to estimate the effect of the dynamical friction in the outer regions of clusters of galaxies (Gambera et al. in preparation) and to compare the observational data with the theoretical ones. This is a good method to test how important is the role of the dynamical friction in the collapse of the clusters of galaxies.

Acknowledgements

We are grateful to V. Antonuccio-Delogu for helpful and stimulating discussions during the period in which this work was performed and the referee Dr. B. S. Ryden for some useful comments.

References

Antonuccio-Delogu, V., 1992, Ph.D. dissertation, ISAS, Trieste (A92)
Antonuccio-Delogu, V., Atrio-Barandela, F., 1992, ApJ., 392, 403 (AA92)
Antonuccio-Delogu, V., Colafrancesco, S., 1994, ApJ., 427, 72 (AC94)
Bardeen, J.M, Bond, J.R., Kaiser, N., Szalay, A.S., 1986, ApJ., 304, 15
Bower R.G., 1991, MNRAS, 248, 332
Chandrasekhar, S., 1943, ApJ., 97, 255
Chandrasekhar, S., von Neumann, J., 1942, ApJ., 95, 489
Chandrasekhar, S., von Neumann, J., 1943, ApJ., 97, 1
Colafrancesco, S., Vittorio, N., 1993, in Cosmic velocity fields, ed. F.R. Bouchet, M. Lachieze-Rey, (Gif-sur-Yvette: Editions Frontieres), 533
Colafrancesco, S., Lucchin, F., Matarrese, S., 1989, ApJ., 345, 3
Colafrancesco, S., Antonuccio-Delogu, V., Del Popolo, A., 1995, ApJ., 455, 32
Croft, R.A.C., Efstathiou, G., 1994, in Numerical simulations in astrophysics: modelling the dynamics of the Universe, ed. J. Franco et al., (Cambridge: Cambridge Univ. Press), 40
Del Popolo, A., Gambera, M., 1997, A. & A. 321, 691 (DG97)
Del Popolo, A., Gambera, M., Antonuccio-Delogu, V., 1996, Mem. Soc. Astr. It., 67, n. 4, 957
Dutta, S.N., Spergel, D.N., 1994, in Numerical simulations in astrophysics: modelling the
dynamics of the Universe, ed. J. Franco et al., (Cambridge: Cambridge Univ. Press),
68
Efstathiou, G., 1994, in Cluster of Galaxies, ed. F. Durret, A. Mazure & J. Tran Thanh
Van, (Gif-sur-Yvette: Editions Frontières), 315
Gambera, M., 1997, Ph.D. dissertation, Istituto di Astronomia dell’ Università di Catania
Gunn, J.E., Gott, J.R., 1972, ApJ., 176, 1 (GG72)
Holtsmark, P.J., 1919, Phys. Z., 20, 162
Kaiser, N., 1993, in Cosmic velocity fields, ed. F.R. Bouchet, M. Lachieze-Rey, (Gif-sur-
Yvette: Editions Frontières), 533
Kandrup, H.E., 1980, Phys. Rep., 63, n. 1, 1 (K80)
Kashlinsky, A., 1986, ApJ., 306, 374
Kolb, E.W., Turner, M.S., 1990, The early Universe (Addison-Wesley)
Liddle, A.R., Lyth, D.H.,1993, Phys. Rep., 231, n 1, 2
Mosconi, M.B., Tissera, P.B., Lambas, D.G., 1994, in Numerical simulations in astrophysics: modelling the dynamics of the Universe, ed. J. Franco et al., (Cambridge:
Cambridge Univ. Press), 63
Peebles, P.J.E., 1993, Principles of Physical Cosmology, Princeton University Press
Ryden, B.S., 1988, ApJ., 329, 589
Ryden, B.S., & Gunn, J.E., 1987, ApJ., 318, 15
Sutherland, W., Dalton, G., 1994, in Cluster of Galaxies, eds. F. Durret, A. Mazure & J.
Tran Thanh Van, (Gif-sur-Yvette: Editions Frontières), 331

This article was processed by the author using Springer-Verlag \LaTeX A&A style file 1990.