THE IDENTIFICATION OF THE OPTICAL COMPANION TO THE BINARY MILLISECOND PULSAR J0610–2100 IN THE GALACTIC FIELD

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ABSTRACT

We have used deep V and R images acquired at the ESO Very Large Telescope to identify the optical companion to the binary PSR J0610–2100, one of the black-widow millisecond pulsars recently detected by the Fermi Gamma-ray Space Telescope in the Galactic plane. We found a faint star (V ∼ 26.7) nearly coincident (δv ∼ 0′′28) with the pulsar nominal position. This star is visible only in half of the available images, while it disappears in the deepest ones (those acquired under the best-seeing conditions), thus indicating that it is variable. Although our observations do not sample the entire orbital period (P = 0.28 days) of the pulsar, we found that the optical modulation of the variable star nicely correlates with the pulsar orbital period and describes a well-defined peak (R ∼ 25.6) at Φ = 0.75, suggesting a modulation due to the pulsar heating. We tentatively conclude that the companion to PSR J0610–2100 is a heavily ablated very low mass star (∼0.02 M⊙) that completely filled its Roche lobe.

Key words: binaries: general – pulsars: individual (PSR J0610-2100) – stars: imaging – techniques: photometric

1. INTRODUCTION

It is generally accepted that millisecond pulsars (MSPs) are formed in binary systems containing a neutron star (NS) that is eventually spun up through mass accretion from an evolving companion. Among these systems those characterized by relatively small eccentricity and very small mass function (typically the companion mass is only Mcom < 0.1 M⊙) are classified as “black-widow” pulsars (BWP). In several cases, the pulsar shows eclipses in the radio signal suggesting that the companion is a non-degenerate, possibly bloated star. In some cases, the eclipse of the radio signal is so extended that it implies a size of the companion larger than its Roche lobe, suggesting that the obscuring material is plasma released by the companion because of the energy injected by the pulsar. However, since the size of the eclipse depends on the inclination angle (King & Beer 2005), not all BWPs are expected to show eclipses. As suggested by King et al. (2003), the formation of BWPs needs two phases: a first one in which the companion spins-up the NS to millisecond periods and a second where the companion is ablated by the pulsar. While it is difficult to describe the two phases using the same star as a companion to the MSP, in globular clusters (GCs), where encounters and exchange interactions are frequent, the white dwarf (WD) companion responsible for the pulsar spinning-up can be replaced by a main-sequence star via an exchange interaction. The following evolution of this newly assembled binary system can cause the progressive vaporization of the companion because of the energy injected by the MSP. Since dynamical interactions are less probable in low-density environments, BWPs were thought to be mainly generated in GCs and then ejected in the field. However, the increasing number of BWPs discovered in the Galactic field suggests that they must form in the disk as well. In this paper, we focus on a BWP in the galactic plane: PSR J0610–2100.

PSR J0610–2100 is an MSP with period P = 3.8 ms and a radio flux at 1.4 GHz of S1.4 = 0.4 ± 0.2 mJy discovered during the Parkes High-Latitude Pulsar Survey (Burgay et al. 2006, hereafter B06). The period derivative ˙P = 1.235 × 10−20 s−1 implies a characteristic age τ = 5 Gyr, a magnetic field B = 2.18 × 108 G, and a rotational energy loss rate ˙E = 2.3 × 1033 erg s−1, similar to the values measured for other MSPs. PSR J0610–2100 is in a binary system with an orbital period of ∼0.28 days. In particular, it is one of 53 binary MSPs with P < 10 ms currently known in the Galactic disk (http://www.atnf.csiro.au/research/pulsar/psrcat/expert.html; Manchester et al. 2005). It is located at a distance of 3.5 ± 1.5 kpc, estimated from its dispersion measure (DM = 60.666 pc cm−3) and the Galactic electron density model of Cordes & Lazio (2002). The pulsar has a proper motion of μ̂cosδ = 7 ± 3 mas yr−1 and μδ = 11 ± 3 mas yr−1 (B06), which implies a transverse velocity of 228 ± 53 km s−1, one of the highest measured for Galactic MSPs. The system mass function (f = 5 × 10−6) implies a lower limit of 0.02 M⊙ for the mass of the companion, assuming 1.35 M⊙ for the pulsar (B06). Thus, in agreement with the definition above, PSR J0610–2100 is probably a BWP seen at a low inclination angle (in fact no eclipse is detected).

Until a few years ago just two other BWPs were known in the Galactic field, namely, PSR B1957+20 (Fruchter et al. 1988b) and PSR J2051–0827 (Stappers et al. 1996a). But very recently, thanks both to dedicated surveys of γ-ray sources and to new blind searches, seven new BWPs have been discovered, most of them having been detected in γ-rays (see Roberts 2011 and references therein). Also, PSR J0610–2100 has been detected in γ-rays by the Large Area Telescope (LAT) on board the Fermi Gamma-ray Space Telescope. Based on positional coincidence with the LAT error box, it was initially associated with the γ-ray source 1FGL J0610.7–6) implies a lower limit of 0.02 M⊙ for the mass of the companion, assuming 1.35 M⊙ for the pulsar (B06). Thus, in agreement with the definition above, PSR J0610–2100 is probably a BWP seen at a low inclination angle (in fact no eclipse is detected).

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Table 1
Summary of the Adopted Data Set and of Detections at Different Orbital Phases of the Companion to PSR J0610-2100

| Filter    | N × texp | Night               | PSR Phase | Detection |
|-----------|----------|---------------------|-----------|-----------|
| V_BESSEL | 6 × 1010 s | 2004 Dec 17         | 0.04–0.26 | NO        |
| V_BESSEL | 3 × 1010 s | 2004 Dec 20         | 0.58–0.67 | YES       |
| R_SPECIAL | 5 × 590 s | 2004 Dec 14         | 0.73–0.83 | YES       |
| R_SPECIAL | 5 × 590 s | 2004 Dec 21         | 0.06–0.16 | NO        |
| R_SPECIAL | 5 × 590 s | 2005 Jan 5          | 0.26–0.36 | NO        |
| R_SPECIAL | 5 × 590 s | 2005 Jan 6          | 0.76–0.86 | YES       |

Note. a Detected in only two images out of three.

Studying the optical emission properties of binary MSP companions is important to better constrain the orbital parameters and to clarify the evolutionary status of these systems and then to track back their history and characteristic timescales. In spite of their importance, only two optical companions to BWP in the Galactic field have been detected to date (Fruchter et al. 1988a; Reynolds et al. 2007; Van Kerkwijk et al. 2011; Stappers et al. 1996b, 1999, 2000). The binary MSP PSR B1957+20 is the first discovered BWP and one of the best-studied members of this class. The optical companion to PSR J1957+20 was identified by Kulkarni et al. (1988), while subsequent observations found the companion to vary by a fraction of 30%–40% in flux over the course of the orbital period (Callanan et al. 1995). Reynolds et al. (2007), modeling the light curve over all the orbital period, constrained the system inclination 63° < i < 67° and the filling factor of the Roche Lobe (0.81 < f < 0.87). Moreover, they ruled out the possibility that the companion is WD, suggesting that it is most likely a brown dwarf. A recent spectroscopic analysis, combined with the knowledge of the inclination angle inferred from models of the light curve, suggested that the PSR B1957+20 is massive with $M_{\text{PSR}}$ = 2.4 $M_\odot$ ($M_{\text{PSR}}$ > 1.66 $M_\odot$ being conservative; van Kerkwijk et al. 2011). The optical companion to binary PSR J2051–0827 was identified by Stappers et al. (1996b). They found that the amplitude of the companion’s light curve was at least 1.2 mag, and that the variation was consistent with the companion’s rotating synchronously about the pulsar and one side being heated by the impinging pulsar flux. In subsequent works it has been possible to study the entire light curve, measuring amplitudes of 3.3 and 1.9 mag in the R band and I band, respectively. The companion star has been modeled by a gravitationally distorted low-mass secondary star which is irradiated by the impinging pulsar wind. The resulting best-fit model is of a Roche lobe filling companion star which converts approximately 30% of the incident pulsar spin down energy into optical flux (Stappers et al. 2000).

Here we present the first identification of the optical companion to PSR J0610–2100, from data acquired at the ESO Very Large Telescope (VLT). The observations and data analysis are described in Section 2, while the results are presented in Section 3 and discussed in Section 4.

2. OBSERVATIONS AND DATA ANALYSIS

The photometric data set used for this work consists of a series of ground-based optical images acquired with the FOcal Reducer/low dispersion Spectrograph 2 (FORS2) mounted at the ESO-VLT. We used the Standard Resolution Collimator, with a pixel scale of 0.25 pixel−1 and a field of view (FOV) of 6.8 × 6.8. In order to obtain deep images free from the blooming due to heavy saturation of bright stars, which can significantly limit the search for faint objects, all the brightest stars in the FOV have been covered with occulting masks.

Six short acquisition images (of 5 s each) and a total of 29 deep images in the V_BESSEL and R_SPECIAL bands (V and R hereafter) were collected during six nights, from mid 2004 December to the beginning of 2005 January (see Table 1), under program 074.D − 0371(A) (PI: E. Sabbi). These data allow us to sample ~25% of the orbital period in V and less than 40% of it in R (see Column 4 in Table 1).

By following standard reduction procedures, we corrected the raw images for bias and flat field. In particular, in order to obtain high-quality master-bias and master-flat images, we selected a large number of images obtained during each observing night and, for each filter, we properly combined them by using the tasks zerocombine and flatcombine in the IRAF package ccddRed. The calibration files thus obtained have been applied to the raw images by using the dedicated task ccddproc.

Based on the word coordinate system of the images, we approximately located the pulsar position and decided to limit the photometric analysis to a region of 500 pixel × 500 pixel (~125° × 125°) centered on it.

We carried out the photometric analysis by using DAOPHOT (Stetson 1987, 1994). We modeled the point-spread function (PSF) in each image by using about 40 bright, isolated, and not saturated stars. The PSF model and its parameters have been chosen by using the DAOPHOT PSF routine on the basis of a χ² test, with a Moffat function (Moffat 1969) providing the best fit to the data in all cases.

Since our purpose is to detect the faintest stars in the field, we selected the three images obtained under the best-seeing conditions (~0.6) in both filters and we combined them using the IRAF task imcombine in order to obtain a high signal to noise master frame to be used as a reference for the object detection. This was done by using the DAOPHOT FIND routine and imposing a detection limit of 3σ.

Finally, by using allframe (Stetson 1987, 1994) we forced the object detection and PSF fitting in each single image adopting the star positions in the master frame as reference. This procedure also allowed us to achieve an improved determination of the star centroids and a better reconstruction of the star intensity profiles.

Photometric calibration. For a straightforward comparison with theoretical models able to provide information about luminosity and temperature of the companion star, we decided to...
calibrate the instrumental magnitudes to the standard Johnson photometric system. To this aim, we first derived the calibration equation for 10 standard stars in the field PG0231 (Stetson 2000), which has been observed with FORS2 during a photometric night (2004 December 17) in both $V$ and $R$ under a calibration program. To analyze the standard star field we used the DAOPHOT PHOT task and performed aperture photometry with a radius $r = 14$ pixels beyond which the contribution of the PSF wings to the star intensity profile becomes negligible. We then compared the obtained magnitudes with the standard Stetson catalog available on the CADC Web site. The resulting calibration equations are $V = v - 0.092(v - r) + 27.79$ and $R = r - 0.019(v - r) + 27.97$, where $v$ and $r$ are the instrumental magnitudes. The color coefficient is very small for the $R$ band, while it could be not negligible in the $V$ band.

**Astrometry.** Since the pulsar timing position is known with a very high precision, obtaining accurate astrometry is a critical requirement in searching for the optical companions to binary MSPs. For this reason, particular care has been devoted to obtain a very good astrometric solution. Since most of the astrometric standard stars are saturated in our catalog, we used the short $R$ band to this purpose. As a first step in our procedure, we registered the pixel coordinates of this image onto the absolute astrometric solution of the FORS2 images.

At the end of the procedure the typical accuracy of the astrometric solution was $\sim 0.2$ arcsec in both right ascension ($\alpha$) and declination ($\delta$).

### 3. The Companion to PSR J0610–2100

In order to identify the companion to PSR J0610–2100 we first searched for objects with coordinates compatible with the nomin al PSR position: $\alpha = 06^h10^m13.59214(10)$ and $\delta = -21^\circ00'28''.0158(17)$ at Epoch MJD = 53100 (B06). Since the epoch of observations is within less than one year from the epoch of the reference radio position, we neglected the effect of proper motion, which is much smaller than the accuracy of the astrometric solution of the FORS2 images.

A first visual inspection of the pulsar region clearly shows that only one star lies within a couple of arcseconds from the MSP radio position: it is located at $\alpha = 06^h10^m13^s15.8$ and $\delta = -21^\circ00'27''.83$, just 0.28 arcsec from PSR J0610–2100. Thus, from positional coincidence alone, we found a very good candidate companion to PSR J0610–2100. Note that the chance coincidence probability that a star is located at the pulsar position is only $P = 0.0007$. Hence, this star is the companion to PSR J0610–2100 with a probability of $\sim 99\%$. Interestingly enough, this star was not present in the master object list obtained from the stacking of the $R$-band images because it is visible in only half of the images, while it completely disappears in the others (see Figure 1). We performed a detailed photometric analysis of this star for measuring its magnitude in as many images as possible, and we found that it is not detected in the deepest $R$ images obtained under the best-seeing conditions (FWHM $\sim 0''.6$). In summary, we were able to measure the magnitude of the star in only 12 images (10 in the $R$ band and 2 in the $V$ band), finding significant variations: $\delta R \sim 1$ mag, from $R = 25.3 \pm 0.1$ to $R = 26.3 \pm 0.2$, and $\delta V \sim 0.5$ mag, from $V = 26.7 \pm 0.2$ to $V = 27.2 \pm 0.2$. In the remaining images, the star magnitude is below the detection threshold ($R = 27 \pm 0.3$ and $V = 27.3 \pm 0.3$), thus suggesting a much more pronounced optical variation. Considering the entire data set, the object’s photometry shows a quite large scatter, significantly ($>5\sigma$) larger than that computed for stars of similar magnitude in the same FOV (see Figure 2). These findings confirm that this is a variable object near the detection limit of our sample.

In order to establish a firm connection between this star and the pulsar we computed the $V$ and $R$ light curves folding each measurement with the orbital period ($P = 0.2860160010$ days) and the ascending node ($T_0 = 52814.249433$) from the radio ephemeris (B06). As shown in Figure 3, although the available data do not allow a complete coverage of the orbital period, the flux modulation of the star nicely correlates with the pulsar orbital phase. The available data are consistent with the rising (in the $V$ band) and the decreasing (in $R$ band) branches of a light curve with a peak at $\Phi = 0.75$. This is the typical

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8 http://www2.cadc-ccda.hia-iha.nrc-cnrc.gc.ca/community/STETSON/standards/

7 CataXcorr is a code aimed at cross-correlating catalogs and finding astrometric solutions, developed by P. Montegriffo at the INAF–Osservatorio Astronomico di Bologna. This package has been used in a large number of papers of our group in the past 10 years.

8 The chance coincidence probability is calculated as $P = 1 - \exp(-\pi \sigma R^2)$ where $\sigma$ is the stellar density of stars with similar magnitude to the candidate companion and $R$ is the accuracy of the astrometric solution.
behavior expected when the surface of the companion is heated by the pulsar flux and the orbital plane has a sufficiently high inclination angle. In fact, in this configuration a light curve with a maximum at $\Phi = 0.75$ (corresponding to the pulsar inferior conjunction, when the companion surface faces the observer) and a minimum at $\Phi = 0.25$ (corresponding to the pulsar superior conjunction) is expected. Indeed the star is not detectable at the epochs corresponding to the orbital phases where the luminosity minimum is predicted.

Based on all these pieces of evidence we propose the identified variable star as the companion to the pulsar; according to previous papers (see Ferraro et al. 2001, 2003; Cocozza et al. 2006; Pallanca et al. 2010), we name it COM J0610−2100.

Since the available $V$ and $R$ measurements are mainly clustered toward the maximum of the emission, but do not allow us to precisely determine it, we used a simple sinusoidal function\(^9\) to obtain a first-guess modeling of the light curve. In the following analysis, we will use these values instead of the mean magnitudes over the entire orbital period. In fact, while the latter would be more appropriate in general, they are not available in this case because of the incomplete sampling of the period. Also note that during the calibration procedure the color term entering the equations has been computed as the difference between the average value of the available $V$ and $R$ instrumental magnitudes. While this is strictly correct for non-variable stars, in the case of COM J0610−2100 it could have introduced an error in the estimated magnitudes. Since in the calibrating equations the coefficients of the color terms are very small, especially in the $R$ band, this uncertainty should be negligible. The resulting magnitudes of COM J0610−2100 at maximum are $R = 25.6$ and $V = 26.7$.

\(^9\) Although this assumption is not supported by a physical reason, it provides a first estimate of the magnitude and color of COM J0610-2100 at maximum.

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\footnote{From NED, NASA/IPAC Extragalactic Database–Galactic Extinction Calculator available at the Web site http://ned.ipac.caltech.edu/forms/calculator.html.}
According to Eggleton (1983), we assume
\[ R_{RL} \simeq \frac{0.49 q^{2/3}}{0.6 q^{2/3} + \ln (1 + q^{1/3})}, \]
where \( q \) is the ratio between the companion and the pulsar masses (\( M_{\text{COM}} \) and \( M_{\text{PSR}} \), respectively). This relation can be combined with the PSR mass function \( f(i, M_{\text{PSR}}, M_{\text{COM}}) = (M_{\text{COM}} \sin i)^3 / (M_{\text{COM}} + M_{\text{PSR}})^2 \) by assuming a nNS mass \( M_{\text{PSR}} = 1.5 M_\odot \) (as recently estimated for recycled pulsars by Ozel et al. 2012; see also Zhang et al. 2011; Kızıltan 2011), thus yielding \( R_{RL}(i) \approx 0.24-0.47 R_\odot \), depending on the inclination angle \( i \) of the orbital system. These values are about 1.7–3.4 times larger than \( R_{BB} \). In the following discussion we assume the value of the Roche lobe as a measure of the size of \( \text{COM} \) J0610–2100 and we discuss how the scenario would change by using \( R_{BB} \) instead of \( R_{RL} \). While these assumptions trace two extreme possibilities, the situation is probably in the midway. In fact, in the case of a completely filled Roche lobe, the mass lost from the companion should produce some detectable signal in the radio band (unless for very small orbital inclinations) and ellipsoidal variations could be revealed in the light curve (unless the heating from the pulsar is dominating).

Under the assumption that the optical variation shown in Figure 3 is mainly due to irradiation from the MSP reprocessed by the surface of \( \text{COM} \) J0610–2100, we can estimate how the re-processing efficiency depends on the inclination angle and, hence, on the companion mass. To this end, we compare the observed flux variation (\( \Delta F_{\text{obs}} \)) between the maximum (\( \Phi = 0.75 \)) and the minimum (\( \Phi = 0.25 \)) of the light curve, with the expected flux variation (\( \Delta F_{\text{exp}} \)) computed from the rotational energy loss rate (\( E \)). Actually, since we do not observe the entire light curve, \( \Delta F_{\text{obs}} \) can just put a lower limit to the re-processing efficiency. Moreover, since these quantities depend on the inclination angle of the system (see below) we can just estimate the reprocessing efficiency as function of \( i \).

At first we have to convert the observed magnitude variation into a flux. We limited our analysis to the \( R \) band since we have more observations and a more reliable sampling of the light curve. At maximum (\( \Phi = 0.75 \)) we assume \( R = 25.6 \), and between \( \Phi = 0.75 \) and \( \Phi = 0.25 \) we estimate an amplitude variation \( \Delta R \gtrsim 1.5 \). Hence we obtained \( \Delta F_{\text{obs}} \sim 1.88 \times 10^{-30} \text{erg s}^{-1} \text{cm}^{-2} \text{Hz}^{-1} \) and considering the filter width (\( \Delta \lambda = 165 \text{nm} \)) we have \( \Delta F_{\text{obs}} \sim 3.4 \times 10^{-15} \text{erg s}^{-1} \text{cm}^{-2} \).

On the other hand, the expected flux variation between \( \Phi = 0.75 \) and \( \Phi = 0.25 \) is given by
\[ \Delta F_{\text{exp}}(i) = \frac{\eta}{A^2} R_{\text{COM}}^2 \frac{1}{4 \pi d_{\text{PSR}}^2} \varepsilon(i), \]
where \( \eta \) is the re-processing efficiency under the assumption of isotropic emission, \( A \) is semimajor axis of orbit which depends on the inclination angle, \( R_{\text{COM}} \) is the radius of the companion star, which we assumed to be equal to \( R_{RL}(i) \), \( d_{\text{PSR}} \) is the distance of pulsar (3.5 kpc)\(^{11} \), and \( \varepsilon(i) \) is the fraction of the re-emitting surface visible to the observer.\(^{12} \) By assuming \( \Delta F_{\text{obs}} = \Delta F_{\text{exp}}(i) \) between \( \Phi = 0.75 \) and \( \Phi = 0.25 \), we can derive a relation

\(^{11} \) In these calculations we adopted a distance of 3.5 kpc, while we discuss below how the scenario changes by varying the distance between the range of values within the quoted uncertainty.

\(^{12} \) In the following, we assume \( \varepsilon(i) = i / 180 \). In fact, for a face-on configuration \( (i \approx 0^\circ) \) no flux variations are expected, while for an edge-on system \( (i = 90^\circ) \) the fraction of the heated surface that is visible to the observer varies between 0.5 (for \( \Phi = 0.75 \)) and 0 (for \( \Phi = 0.25 \)).

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Figure 5. Reprocessing efficiency for isotropic emission (\( \eta \)) as a function of the inclination angle \( i \) calculated assuming \( M_{\text{PSR}} = 1.5 M_\odot \) and \( R_{COM} = R_{RL}(i) \). The corresponding values for the companion mass in units of \( M_\odot \) are reported on the top axis. The solid line corresponds to the lower limit of the amplitude of variation (\( \Delta R = 1.5 \)), while the gray region to values of \( \Delta R \) up to 3, that should be appropriate if considering the entire light curve (i.e., see PSR J2051–0827; Stappers et al. 1996b). The shaded area marks the region of the diagram where an anisotropic emission of the re-processed flux is admitted (in case of isotropy, in fact, the efficiency would be unphysical; \( \eta > 100\% \)). The dotted line marks the physical limit for core hydrogen-burning stars (i.e., objects with masses \( \geq 0.08 M_\odot \)).

 linking the re-processing efficiency and the inclination angle. The result is shown in Figure 5. The absence of eclipses in the radio signal allows us to exclude very high inclination angles. As shown in Figure 5, pulsar companions with stellar mass above the physical limit for core hydrogen-burning star (i.e., with \( M \geq 0.08 M_\odot \)) necessarily imply a non-isotropic emission mechanism of the pulsar flux (otherwise a larger than 100% physical efficiency would be required). On the contrary, a re-processing efficiency between 40% and 100% is sufficient for less massive companions and intermediate inclination angles.

Taking into account the uncertainty on the pulsar distance, only re-processing efficiencies larger than ~60% for inclination angles in excess to 50° and companion masses lower than ~0.03 \( M_\odot \) are allowed in the case of the distance upper limit (5 kpc). Instead, in case of a closer distance \( \eta \) decreases for all inclination angles, thus making companion stars with masses larger than 0.08 \( M_\odot \) also acceptable. For instance, for companions masses between 0.08 and 0.2 \( M_\odot \) and a 2 kpc distant pulsar, the re-processing efficiency ranges between 30% and 60% for any value of \( i \).

The observed optical modulation can be reproduced considering a system seen at an inclination angle of about 60°, with a very low mass companion (\( M_{\text{COM}} \approx 0.02 M_\odot \)) that has filled its Roche lobe, and a re-processing efficiency of about 50%. On the other hand, if we use \( R_{BB} \) instead of \( R_{RL} \), the efficiency becomes larger than 100% for every inclination angle, and the only possible scenario would be that of an anisotropic pulsar emission. However, even with this assumption it is very difficult to obtain an acceptable value for \( \eta \). This seems to confirm that \( R_{BB} \) is too small to provide a good estimate of the star physical size. Forthcoming studies will allow us to better constrain the system parameters.
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