The XMM-Newton/2dF survey - V. the radio properties of the X-ray population

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ABSTRACT
In this paper we cross-correlate the FIRST 1.4 GHz radio survey with a wide field (1.6 deg²) shallow \([f_X(0.5 – 8 \text{ keV}) \approx 10^{-14}\text{ erg s}^{-1}\text{ cm}^{-2}]\) XMM-Newton survey. We find 12 X-ray/radio matches representing 4% of the X-ray selected sample. Most of them are found to be associated with AGNs (total of 9) on the basis of the observed optical spectra (3), radio morphology (2) or X-ray/optical properties (4), while one radio source is identified with an X-ray selected cluster. We also find two sources associated with low redshift galaxies with narrow emission line optical spectra, X-ray luminosity \(L_X(0.5 – 8\text{ keV}) \approx 10^{41}\text{ erg s}^{-1}\), radio luminosity density \(L_{1.4\text{ GHz}} \approx 5 \times 10^{22}\text{ W Hz}^{-1}\) and \(\log f_X / f_{\text{opt}} \approx -2\) suggesting ‘normal’ star-forming galaxies. We argue that radio surveys combined with X-ray samples could provide a powerful tool for identifying X-ray selected ‘normal’ galaxies powered by stellar processes. Finally, radio loud and quiet systems in the present sample have mean X-ray spectral properties consistent with \(\Gamma \approx 1.9\).

Key words: Surveys – Galaxies: active – X-rays: galaxies – X-rays: general

1 INTRODUCTION
X-ray surveys are not only dominated by AGNs but are also believed to be the least biased method for selecting such systems (e.g. Brinkmann et al. 2000). The evidence above has motivated a number of studies aiming to compile X-ray selected (and thus least biased) AGN samples to understand the large variations in their observational properties and to test unification schemes. At radio wavelengths in particular large programs have been initiated by cross-correlating wide area radio surveys with archival X-ray data aiming to understand the radio loud/quiet AGN dichotomy (Brinkmann et al. 1997; Caccianiga et al. 1999; Brinkmann et al. 2000).

For example Brinkmann et al. (2000) cross-correlated the FIRST radio survey (Faint Images of the Radio Sky at Twenty centimeters; Becker et al. 1995) with the ROSAT All Sky Survey (RASS-II; Voges et al. 1999) at a 0.1-2.4 keV flux limit of \(\approx 10^{-13}\text{ erg s}^{-1}\text{ cm}^{-2}\). About 1/3 of the X-ray sources are identified with radio counterparts. The above fraction of X-ray/radio associations is higher than that found in previous soft X-ray selected samples (Stocke et al. 1991; Zamorani et al. 1999; Ciliegi et al. 1995; see Ciliegi et al. 2003 for a summary), most likely due to the relative X-ray and radio limits of the RASS-II and the FIRST surveys. Follow-up optical spectroscopy of the X-ray/radio matches shows that the majority of these sources are AGNs spanning both the radio loud and the radio quiet regime. These data indicate that there is no sharp boundary between radio loud and radio quiet AGNs but rather a smooth transition. A highly significant X-ray–radio luminosity correlation for these sources is also reported. Apart from AGNs the sample of Brinkmann et al. (2000) also comprises a small fraction (≈ 5%) of ‘normal’ galaxies likely to be starbursts.

At fainter X-ray flux limits (≈ \(10^{-16}\text{ erg s}^{-1}\text{ cm}^{-2}\)) Bauer et al. (2002) investigated the radio properties of the X-ray sources using data from the 1 Ms Chandra Deep Field North (CDF-N) combined with \(\mu\text{Jy}\) sensitivity VLA observations. They find that \(\approx 38 \pm 20\%\) of the X-ray selected AGNs have radio counterparts to the limit of the VLA data (0.05 mJy). They also argue that the radio emission, at least in the sub-sample of harder (i.e. obscured) AGNs, is likely
to be associated with circumnuclear star-formation activity. Contrary to AGNs the largest overlap between X-ray and radio sources (66 ± 25%) is for the sub-sample of narrow emission line galaxies likely to be star-forming systems. The evidence above suggests the emergence of ‘normal’ and starburst galaxies at faint X-ray fluxes.

In addition to soft X-ray selected samples a number of hard X-ray surveys have been followed by radio observations yielding X-ray/radio matched samples with identification rates $\gtrsim 30\%$ (e.g. Ciliegi et al. 2003; Akiyama et al. 2000; Berger et al. 2001; Georgakakis et al. 2004b). Ciliegi et al. (2003) argue that this is likely due to observational effects (e.g. deeper radio data) and the fact that both the radio wavelengths and the hard X-ray energies ($> 2\text{keV}$) are least biased to obscuration effects.

The studies above mainly concentrate on either very bright ($> 10^{-13}\text{erg s}^{-1}\text{cm}^{-2}$) or extremely faint ($< 10^{-15}\text{erg s}^{-1}\text{cm}^{-2}$) X-ray sources. In the present study we explore the radio properties of the X-ray population in the intermediate X-ray flux range ($\approx 10^{-14}\text{erg s}^{-1}\text{cm}^{-2}$). We cross correlate the FIRST radio survey with a wide area (1.6 deg$^2$) shallow $[f_X(0.5 - 8\text{keV}) \approx 10^{-14}\text{erg s}^{-1}\text{cm}^{-2}]$ XMM-Newton survey near the North Galactic Pole region (XMM-Newton/2dF survey†). Section 2 describes the X-ray, radio and optical data used in the present study while section 3 presents the X-ray/radio matched sample. The results are discussed in section 4 with section 5 summarising our conclusions. Throughout this paper we adopt $H_0 = 65\text{km s}^{-1}\text{Mpc}^{-1}$, $\Omega_M = 0.3$ and $\Omega_\Lambda = 0.7$.

2 OBSERVATIONS

2.1 X-ray data

The X-ray sample used in the present study is compiled from the XMM-Newton/2dF survey. This is a wide area ($\approx 2.5\text{deg}^2$) shallow $[f_X(0.5 - 8\text{keV}) \approx 10^{-14}\text{erg s}^{-1}\text{cm}^{-2}]$, 5$\sigma$ survey carried out by the XMM-Newton near the North and South Galactic Pole regions. The data reduction, source extraction, flux estimation and catalogue generation are described in detail by Georgakakis et al. (2003a, Paper I; 2004a, Paper II). In the present study we concentrate on the North Galactic Pole F864 subregion of the XMM-Newton/2dF survey covering an area of 1.6 deg$^2$. This is because of the wealth of follow-up observations (optical photometry and spectroscopy, radio) available for these fields. The X-ray sample comprises a total of 291 sources detected in the 0.5-8keV spectral band above the 5$\sigma$ threshold. We note that about 10% of the surveyed area is covered at the flux $f_X(0.5 - 8\text{keV}) = 10^{-14}\text{erg s}^{-1}\text{cm}^{-2}$. This fraction increases to about 50% at $f_X(0.5 - 8\text{keV}) = 2 \times 10^{-14}\text{erg s}^{-1}\text{cm}^{-2}$.

2.2 Radio data

The radio data are from the FIRST (Faint Images of the Sky at Twenty centimeters) radio survey (Becker et al. 1995; White et al. 1997). The observations are carried out at 1.4 GHz using the NRAO Very Large Array (VLA) in the B-configuration. The limiting flux density is $S_{1.4\text{GHz}} = 0.8\text{mJy}$ with a 5$\sigma$ source detection limit of $\approx 1\text{mJy}$. The catalogue is estimated to be 95 and 80% complete at 2 and 1 mJy respectively (Becker et al. 1995). The positional accuracy of the detected sources is better than $\approx 1\text{arcsec}$ facilitating their optical identification.

The radio data used in the present study are from the April 11, 2003 version of the FIRST VLA catalogue. A total of 78 radio sources overlap with the XMM-Newton pointings of the F864 region.

2.3 Optical photometric and spectroscopic data

The XMM-Newton/2dF survey F864 region overlaps with the Sloan Digital Sky Survey (York et al. 2000). The SDSS is an on-going imaging and spectroscopic survey that aims to cover about 10000 deg$^2$ of the sky. Photometry is performed in 5 bands ($ugriz$; Fukugita et al. 1996; Stoughton et al. 2002) to the limiting magnitude $g \approx 23\text{mag}$, providing a uniform and homogeneous multi-color photometric catalogue. The SDSS spectroscopic observations will obtain spectra for over 1 million objects, including galaxies brighter than $r = 17.7\text{mag}$, luminous red galaxies to $z \approx 0.45$ and colour selected QSOs (York et al. 2000; Stoughton et al. 2002). In the present study we use data from the Early Data Release (EDR; Stoughton et al. 2002).

In addition to the SDSS the F864 region overlaps with the recently completed 2dF Galaxy Redshift Survey (2dFGRS‡; Colless et al. 2001; Colless et al. 2003) and the 2dF QSO Redshift Survey (2QZ§; Croom et al. 2001). Both the 2dFGRS and 2QZ are large-scale spectroscopic campaigns that fully exploit the capabilities of the 2dF multi-fibre spectrograph on the 4m Anglo-Australian Telescope (AAT). These projects provide high quality spectra, redshifts and spectral classifications for 220000 $b_j < 19.4\text{mag}$ galaxies and 23000 optically selected $b_j < 20.85\text{mag}$ QSOs.

3 X-RAY – RADIO CORRELATION

In the present study we use a total of 291 XMM-Newton/2dF survey sources with detection significance $> 5\sigma$ in the 0.5-8keV band. The limiting flux of these observations in the 0.5-8keV band is $f_X(0.5 - 8\text{keV}) \approx 10^{-14}\text{erg s}^{-1}\text{cm}^{-2}$.

The FIRST and the XMM-Newton/2dF source catalogues are cross correlated using an initial matching radius of 10 arcsec. This is to include in our sample well separated double lobe radio sources that may be missed if a smaller radius is employed. The candidate identifications were then visually inspected rejecting single component sources (i.e. not double lobes) that lie more than 7 arcsec from an X-ray detection. The final sample comprises a total of 12 X-ray sources with radio counterparts. Given the surface density of the FIRST survey ($\approx 90\text{deg}^2$; Becker et al. 1995) and the XMM-Newton/2dF sources we estimate about 0.3 spurious identifications for the above sample within the 7 arcsec

† http://www.astro.noa.gr/~xray/

‡ http://msowww.anu.edu.au/2dFGRS/

§ http://www.2dfquasar.org
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radius. We note that increasing the initial matching radius to 40 arcsec does not include any additional sources in the sample.

The SDSS is used to optically identify these sources by estimating the probability a given candidate is chance coincidence (Downes et al. 1986). The more accurate radio positions are used for the optical identifications. Of the 12 X-ray/radio matches we propose 10 candidate optical identifications. Optical images of these sources are shown in Figure 1. The optical spectrum of source #4 (see Table 1) is plotted in Figure 2. The optical spectra of the remaining sources are either presented in Paper II or can be found in the 2QZ, 2dFGRS and SDSS archives.

The X-ray/radio matched sample is presented in Table 1 which has the following format:

1. source identification number.

2. Name of the FIRST radio source closest to the X-ray source.

3-4. Right ascension ($\alpha_{1.4GHz}$) and declination ($\delta_{1.4GHz}$) of the radio source position in J2000. In the case of double lobes we list the position of the centroid.

5. Offset in arcseconds between the X-ray source centroid estimated by the Ewavelet task of SAS and the radio source centre. We note that the X-ray source centroid does not always coincide with the peak of the X-ray emission.

6. Offset in arcseconds between the radio source position and the optical source centre.

7. Probability, $P$, the optical identification or a radio source is a chance coincidence.

8. Optical $g$-band magnitude.

9. 1.4 GHz radio flux density.

10. Redshift. The source from which the redshift estimate was obtained is also listed in Table 1.

11. Classification on the basis of the observed optical spectral features or radio morphology: AB: absorption lines only, NL: narrow emission lines, EA: both absorption and emission lines, BL: broad emission lines and DOUBLE: for radio sources with double lobes. We note that source #1

### Table 1. X-ray/radio matched sources in the XMM-Newton/2dF survey.

| ID    | NAME                | $\alpha_{1.4GHz}$ (J2000) | $\delta_{1.4GHz}$ (J2000) | $\Delta\theta_{RX}$ (arcsec) | $\Delta\theta_{RO}$ (arcsec) | $P_{RO}$ (%) | $g$ | $S_{1.4GHz}$ (mJy) | $z$ | Class$^a$ |
|-------|---------------------|---------------------------|---------------------------|-------------------------------|-----------------------------|--------------|-----|-------------------|-----|-----------|
| 1$^b$ | FIRST J134304.6-000055 | 13 43 04.61 | 00 00 55.35 | 3.3 | 0.9 | 0.27 | 21.97 | 28.26 | – | DOUBLE |
| 2     | FIRST J134233.9-001553 | 13 42 33.04 | 00 15 53.24 | 3.6 | 2.9 | 0.44 | 19.98 | 0.73 | 2.132$^3$ | BL |
| 3     | FIRST J134212.2-001737 | 13 42 12.21 | 00 17 37.82 | 4.3 | 0.1 | $<0.01$ | 17.18 | 1.50 | 0.087$^2$ | EA |
| 4     | FIRST J134347.5+002024 | 13 43 47.47 | 00 20 24.22 | 2.9 | 0.5 | 0.01 | 19.55 | 3.13 | 0.240$^1$ | AB |
| 5     | FIRST J134414.2+001642 | 13 44 14.18 | 00 16 42.31 | 1.1 | 0.4 | 0.01 | 19.67 | 95.96 | – | DOUBLE |
| 6     | FIRST J134232.4-003151 | 13 42 32.38 | 00 31 51.10 | 0.4 | 0.1 | $<0.01$ | 18.98 | 1.55 | 1.209$^4$ | BL |
| 7     | FIRST J134133.4-002432 | 13 41 33.36 | 00 24 32.26 | 2.8 | 0.3 | $<0.01$ | 17.08 | 2.98 | 0.072$^2$ | NL |
| 8     | FIRST J134133.7-002555 | 13 41 37.68 | 00 25 55.02 | 2.7 | 0.3 | $<0.01$ | 16.56 | 2.37 | 0.052$^2$ | NL |
| 9     | FIRST J134412.9-003006 | 13 44 12.89 | 00 30 05.51 | 2.5 | 0.1 | $<0.01$ | 20.55 | 8.03 | 0.708$^4$ | BL |
| 10    | FIRST J134351.8-002832 | 13 44 31.77 | 00 28 31.48 | 2.9 | 0.6 | 0.50 | 23.80 | 1.43 | – | – |
| 11    | FIRST J134218.4-003120 | 13 41 28.36 | 00 31 20.21 | 0.2 | 0.4 | 0.04 | 21.55 | 0.81 | – | – |
| 12    | FIRST J134447.0-003009 | 13 44 46.97 | 00 30 09.29 | 0.7 | 0.7 | 0.22 | 22.30 | 50.46 | – | DOUBLE |

$^a$AB: absorption lines; NL: Narrow emission lines; BL: Broad optical emission lines; EA: absorption+emission lines

DOUBLE: double lobe morphology

$^b$this double lobe radio sources is associated with an X-ray cluster.

1: XMM-Newton/2dF spectroscopic program; 2: 2dFGRS; 3: 2QZ; 4: SDSS; 5: Terlevich et al. (1991)

### Table 2. X-ray properties of the X-ray radio matched sources in the XMM-Newton/2dF survey.

| ID | net counts (0.5–8.0 keV) | $f_X$ (0.5–8.0 keV) | $L_X$ (0.5–8.0 keV) | $L_{1.4GHz}$ (W Hz$^{-1}$) | $N_H$ (10$^{20}$ cm$^{-2}$) | $\Gamma$ |
|----|--------------------------|---------------------|---------------------|-----------------------------|----------------------------|--------|
| 1  | 45                       | 4.36 ± 0.98         | –                   | –                           | –                           | –      |
| 2  | 24                       | 2.78 ± 1.03         | (8.74 ± 3.24) × 10$^{44}$ | 7.19 × 10$^{25}$ | < 7.5                     | 1.81$^{+0.35}$– $0.32$ |
| 3  | 21                       | 3.64 ± 1.08         | (7.89 ± 2.34) × 10$^{41}$ | 3.53 × 10$^{22}$ | < 3.0                     | 2.82$^{+1.15}$– $1.05$ |
| 4$^b$ | 141                  | 26.9 ± 2.34         | (5.21 ± 0.45) × 10$^{43}$ | 7.52 × 10$^{21}$ | < 0.6                     | 2.03$^{+0.16}$– $0.23$ |
| 5$^b$ | 766                  | 93.8 ± 3.82         | –                   | –                           | < 1.4                     | 1.76$^{+0.13}$– $0.19$ |
| 6  | 93                       | 11.10 ± 1.45        | (9.25 ± 1.21) × 10$^{44}$ | 2.85 × 10$^{25}$ | < 1.4                     | 2.60$^{+0.41}$– $0.37$ |
| 7  | 22                       | 1.31 ± 0.43         | (1.90 ± 0.63) × 10$^{41}$ | 4.64 × 10$^{22}$ | < 4                      | 2.47$^{+0.99}$– $0.78$ |
| 8  | 30                       | 1.95 ± 0.52         | (1.44 ± 0.39) × 10$^{41}$ | 1.84 × 10$^{22}$ | < 8.7                     | 2.14$^{+0.70}$– $0.23$ |
| 9  | 77                       | 6.11 ± 0.8          | (1.44 ± 0.19) × 10$^{43}$ | 3.23 × 10$^{25}$ | < 0.3                     | 2.80$^{+0.22}$– $0.28$ |
| 10 | 31                       | 1.67 ± 0.40         | –                   | –                           | < 4                       | 2.27$^{+0.48}$– $0.28$ |
| 11 | 110                      | 9.09 ± 0.96         | –                   | –                           | 150$^{+50}$– $50$         | 0.51$^{+0.25}$– $0.22$ |
| 12 | 32                       | 1.95 ± 0.45         | –                   | –                           | 56$^{+61}$– $54$          | 0.88$^{+0.42}$– $0.42$ |

$^a$X-ray flux and X-ray spectral properties are estimated from the MOS

$^b$X-ray spectral fitting used $\chi^2$ statistic with both $\Gamma$ and $N_H$ free parameters
has double lobe morphology and is associated with an X-ray selected cluster.

To explore the X-ray spectral properties of the X-ray/radio matched population we use the C-statistic technique (Cash et al. 1979) as implemented in xspec v11.2. This method has been developed to extract information from X-ray spectra with small number of counts. The data are grouped to have a minimum of one count per bin but using higher binning factors does not modify the results. Assuming an absorbed power law model we attempt to constrain the N\(_H\) fixing the power law index to Γ = 1.9. Only two of the sources presented here (sources #11 and #12) show evidence for absorbed columns > 5 × 10\(^{21}\) cm\(^{-2}\). The results are shown in column 5 of Table 2. We also attempt to constrain the power law index Γ keeping the column density fixed to the Galactic value. The results are also presented in Table 2. With the exception of sources #11 and #12 the present sample has steep X-ray spectra. Source #5 in Tables 1 and 2 has sufficient counts that allow standard χ\(^2\) spectral analysis. This source is also fit by an absorbed power law model with both the column density, N\(_H\), and the X-ray spectral index, Γ, free parameters.

Table 2 presents the X-ray properties of the present sample. We list:

1. source identification number.
2. 0.5-8 keV net source counts.
3. 0.5-8 keV X-ray flux in erg s\(^{-1}\) cm\(^{-2}\).
4. 0.5-8 keV X-ray luminosity in erg s\(^{-1}\), if a redshift is available.
5. 1.4 GHz radio luminosity density in W Hz\(^{-1}\), if a redshift is available.
6. column density N\(_H\) estimated by the C-statistic for a fixed power law index Γ = 1.9. The only exception is source #5 which is fit by an absorbed power law using χ\(^2\) statistics.
7. power law spectral index Γ estimated by the C-statistic for a fixed column density N\(_H\) = 2 × 10\(^{20}\) cm\(^{-2}\). Source #5 is fit using χ\(^2\) statistics.

The sources in Tables 1 and 2 span both the radio loud (#1, 4, 5, 9, 10, 11, 12) and the radio quiet (#2, 3, 6, 7, 8) regime with the radio-loudness parameter defined as in Stocke et al. (1991; see Figure 5 below). In Table 2 both radio quiet and loud systems have X-ray spectral index consistent on average with Γ = 1.9. Coadding the X-ray spectra of radio loud and quiet sources separately yields Γ = 1.71\(^{+0.14}_{-0.14}\) and Γ = 2.19\(^{+0.19}_{-0.31}\) respectively. Although radio loud sources are flatter the spectral indices above are consistent within the 90% confidence level.

Figure 3 plots g-band magnitude against 0.5-8 keV X-ray flux for both X-ray sources with radio counterparts and the whole X-ray selected sample. The log (f\(_X\)/f\(_{opt}\)) = ±1 lines in this figure delineate the region of the parameter space occupied by AGNs. X-ray sources with radio emission span a wide range of X-ray–to–optical flux ratios. Although many X-ray/radio matches lie in the AGN region of the parameter space, a small number of sources have log f\(_X\)/f\(_{opt}\) < −1 suggesting relatively optically luminous obscured AGN and Low Luminosity AGNs (LLAGN; Lehmann et al. 2001) or ‘normal’ galaxies.

The lines of constant X-ray–to–optical flux ratio in Figure 3 are estimated from the 0.5-8 keV flux, f\(_X\)(0.5 – 8 keV), and the g-band magnitude according to the relation

\[
\log \frac{f_X}{f_{opt}} = \log f_X(0.5\text{--}8\text{keV}) + 0.4g + 5.07. \tag{1}
\]

The equation above is derived from the X-ray–to–optical flux ratio definition of Stocke et al. (1991) that involved 0.3–3.5 keV flux and V-band magnitude. These quantities are converted to 0.5-8 keV flux and g-band magnitude assuming a mean colour B−V = 0.8 using the colour transformations of Fukugita et al. (1996) and a power law X-ray spectral energy distribution with Γ = 1.8.

Figure 4 plots total band X-ray luminosity against radio luminosity density for those sources in the present sample with available redshift information. Both L\(_X\) and L\(_{1.4}\) are k-corrected assuming a power law spectral energy distribution with Γ = 1.8 and \(\alpha = 0.8\) respectively. Also shown is the best fit \(L_X - L_{1.4}\) relation for local star-forming galaxies derived by Ranalli, Comastri & Setti (2003). Two of our sample sources (#7 and #8) have \(L_X(0.5\text{--}8\text{keV}) \approx 10^{43}\) erg s\(^{-1}\) cm\(^{-2}\) and lie close to the Ranalli et al. (2003) relation suggesting they are likely to be dominated by starformation. These same sources also have low X-ray–to–optical flux ratios in Figure 3 (log f\(_X\)/f\(_{opt}\) ≈ −2) which is additional evidence for star-formation activity in these systems. The optical spectra for these two sources are presented in Paper II. The remaining sources in Figure 4 either deviate from the \(L_X - L_{1.4}\) relation of local star-forming systems (suggesting the presence of AGN) or have X-ray luminosities well in excess of typical starburst galaxies (\(\approx 10^{42}\) erg s\(^{-1}\)) (Moran, Lehnert & Helfand 1999).

We also find a double lobe radio source (#1) associated with diffuse X-ray cluster emission. As discussed in Appendix A this cluster is likely to lie at \(z \approx 0.6\). Powerful radio galaxies are known to be good tracers of dense

Figure 2. Optical spectrum of source #4 showing absorption lines only at \(z = 0.240\).
environments with Fanaroff-Rilley I type sources (FRI) associated with rich groups or clusters and Fanaroff-Rilley II type sources (FRII) avoiding rich clusters at low-z (e.g. Zirbel 1997 and references therein). Recently, Zanichelli et al. (2001) used NVSS data to compile an intermediate redshift ($z \approx 0.1 - 0.3$) cluster sample based on radio selection. Follow-up optical photometry and spectroscopy confirmed that a significant fraction of their targeted cluster candidates are real, demonstrating the strength of radio selection in cluster studies. Our X-ray/radio matched sample comprises a total of 3 double lobe radio sources (#1, #5 and #12) of which only one (#1) is clearly associated with X-ray cluster emission. The remaining two lie in either poor or high-z clusters/groups that remain undetected in both the X-ray and the optical data. Indeed, as discussed in Appendix A we do not find evidence for optical galaxy overdensity in the vicinity of sources #5 and #12, at least to the optical magnitude limit of the SDSS.

Sources with absorption line optical spectra and high X-ray luminosities ($L_X \approx 10^{42} - 10^{43}$ erg s$^{-1}$), similar to those found in the present sample (#3, 4 in Table 1), have also been identified in previous X-ray surveys (e.g. Griffiths et al. 1995; Blair, Georgantopoulos & Stewart 1997; Allen, Di Matteo & Fabian 2000; Comastri et al. 2002; Brusa et al. 2003). Possible scenarios for the observed X-ray emission in these optically normal (e.g. no emission lines) systems include (i) BL-Lac type activity (e.g. Blair et al. 1997), (ii) Advection Dominated Accretion Flows (ADAF; Narayan & Yi 1995; Di Matteo et al. 2000), (iii) diffuse cluster of group X-ray emission and (iv) LLAGN with its optical signature diluted by the host galaxy light (e.g. Moran, Filippenko & Chornock 2002; Severgnini et al. 2003).

For sources #3, 4 we can exclude the possibility of X-ray cluster emission. The X-ray sources associated with these two galaxies are not extended. Also, no optical galaxy overdensity is found in the vicinity of the two sources (Basillakos et al. 2004, Paper III; Goto et al. 2002). The evidence above is against the cluster emission scenario although the presence of hot gas associated with a group (that remains undetected in both the X-ray and the optical data) cannot be ruled out. Below we argue that source #3 is likely to be associated with an ADAF or a LLAGN, while source #4 is a BL-Lac candidate.

We explore the nature of these sources using their two-point spectral indices, optical-to-X-ray ($\alpha_{OX}$) and radio-to-optical ($\alpha_{RO}$) defined by Stocke et al. (1991):

![Figure 1. SDSS optical r-band images of the X-ray/radio matched sources in Table 1 with the radio contours overlaid. The position of the X-ray centroid is shown with the circle with a radius of 5 arcsec.](image)
\[ \alpha_{OX} = -\log(f_{\text{2keV}}/f_{\text{2500}})/2.605, \]
\[ \alpha_{RO} = \log(S_{\text{5GHz}}/f_{\text{2500}})/5.38, \]

where \( f_{\text{2keV}}, S_{\text{5GHz}}, \) and \( f_{\text{2500}} \) are the monochromatic fluxes at 2 keV, 5 GHz and 2500 Å respectively. The \( S_{\text{5GHz}} \) is estimated from the 1.4 GHz flux density assuming a power law spectral energy distribution of the form \( f_{\nu} \propto \nu^{-\alpha} \) with \( \alpha = 0.8 \). The X-ray flux density at 2 keV is estimated from the 0.5-8 keV flux assuming a mean photon spectral index \( \Gamma = 1.8 \). The monochromatic optical flux at 2500 Å, \( f_{2500} \), is estimated from the SDSS \( g \)-band magnitudes assuming an optical power law spectral slope \( \alpha = 0.5 \) (Brinkmann et al. 2000). The flux densities are not k-corrected but this effect is small and does not modify the estimated monochromatic flux ratios.

The results are shown in Figure 5 plotting \( \alpha_{OX} \) against \( \alpha_{RO} \). Also shown are the regions of the parameter space occupied by different classes of objects as defined by Stocke et al. (1991; see their Figure 6b). The absorption line galaxy in our sample (#4) has \( \alpha_{OX} \) and \( \alpha_{RO} \) consistent with those of BL-Lacs. This coupled with the soft X-ray spectral properties of this system and the absorption line optical spectrum strongly support the BL-Lac scenario. Fossati et al. (1998) showed that the peak frequency of the BL-Lac SED is a function of the radio luminosity of these objects: less luminous systems peak at UV and soft X-ray wavelengths (High energy peak BL-Lacs), while more luminous radio sources have a peak in the infrared regime (Low energy peak BL-Lacs). Source #4 at \( z = 0.240 \) has a radio luminosity at 5 GHz of \( \log \nu L_{\nu} \approx 10^{22} \) erg s\(^{-1} \) suggesting that this is a High energy peak BL-Lac (see Figure 7a of Fossati et al. 1998). The estimated radio luminosity is also consistent with the steep X-ray spectral index of this source in Table 2 (see Figure 11 of Fossati et al. 1998). We also note that absorption line systems with properties similar to those of source #4 have been identified in previous X-ray surveys with available radio observations (Gunn et al. 2003; Brusa et al. 2003).

Three sources in Figure 5 lie close to the ‘normal’ galaxy region of the Stocke et al. (1991) classification scheme. Two of these sources are the ‘normal’ galaxy candidates showing narrow emission line optical spectra. The third source is #3 in Tables 1 and 2 exhibiting both narrow emission and absorption optical lines. The evidence above combined with the fact that this system deviates from the Ranalli et al. (2003) \( L_X - L_{1.4} \) relation in Figure 4 suggest an ADAF or a LLAGN.

Broad optical emission line systems and double lobe sources in Figure 5 occupy the radio loud and the radio quiet AGN space. We note that the one double lobe source associated with the X-ray cluster is not plotted here since its X-ray emission is not due to the central AGN. One of the two spectroscopically unclassified sources (#11) in the present sample has \( \alpha_{OX} \) and \( \alpha_{RO} \) consistent with those of BL-Lacs. This is inconsistent however, with the relatively hard X-ray spectral properties of this source in Table 2 suggesting \( N_H \approx 1.5 \times 10^{22} \) for \( \Gamma = 1.9 \). The second spectroscopically unclassified source (#10) lies in the border line between radio loud QSOs and BL-Lacs.

A detailed description of individual sources is presented in Appendix A. The classification of the present sample into different classes is performed on the basis of their optical spectroscopic (i.e. spectral features), photometric (i.e. resolved or point-like sources) and X-ray properties (i.e. X-ray luminosity, X-ray–to–optical flux ratio, X-ray spectral properties). More information about the source classification can be found in Appendix A. The present sample of X-ray/radio matches comprises: (i) 3 broad emission line AGNs (#2, 6, 9), (ii) 2 sources with double lobe radio morphology also indicating AGN activity (#5, 12), (iii) 1 double lobe source (#1) associated with X-ray cluster emission, (iv) 2 systems with optical spectra dominated by the host galaxy (i.e. absorption and/or narrow emission lines) and X-ray/optical properties suggesting AGN activity (#3, 4), (v) 2 ‘normal’ galaxy candidates exhibiting narrow emission line optical spectra (#7, 8) and (vi) 2 unclassified sources with no optical spectroscopic information (#10, 11). The X-ray/optical properties of this latter class of sources suggests AGN activity. We also note that although the X-ray and optical properties of ‘normal’ galaxy candidates are consistent with stellar origin for the X-ray emission we cannot exclude the possibility of heavily obscured AGN or LLAGN. Summarising the classification above, 9 out of the 12 X-ray/radio matches are associated with AGN activity on the basis of the broad optical emission lines (total of 3), radio morphology (2) or X-ray/optical properties (4). One radio source is associated with X-ray cluster emission with the remaining two sources being ‘normal’ galaxy candidates.
4 DISCUSSION

Brinkmann et al. (2000) cross-correlated the FIRST radio survey with the RASS-II catalogue with a 0.1–2.4 keV limiting flux of \( \approx 10^{-13} \text{ erg s}^{-1} \text{ cm}^{-2} \). These authors found that \( \approx 30\% \) of the RASS-II sources have radio counterparts to the limit of the FIRST survey. In the present survey probing X-ray fluxes about 1 dex fainter than the RASS-II we find that 12 out of 291 X-ray sources (\( \approx 4\% \)) have radio counterparts. This fraction is much lower than that of the RASS-II (\( \approx 30\% \)) suggesting that a significant fraction of the X-ray sources have radio emission below the FIRST flux density limit. Ciliegi et al. (2003) suggest that the fraction of the X-ray/radio matches strongly depends on the relative depths of the X-ray and the radio data. They quantify this effect using the ratio of the X-ray–to–radio flux density limits of different surveys, \( S_{1.4 \text{ GHz}} / f_{2\text{ keV}} \). For the XMM-Newton/2dF survey we estimate \( S_{1.4 \text{ GHz}} / f_{2\text{ keV}} \approx 10^5 \) (assuming \( f = 1.8 \)) which is 1 dex higher than that of \( \approx 10^4 \) for the Brinkmann et al. (2000) sample. Ciliegi et al. (2003) also provide a comparison of the X-ray/radio identification rate of various surveys spanning a range of X-ray and radio limits (see their Table 2). The XMM-Newton/2dF survey has the lowest fraction of X-ray/radio matched sources in this table (\( 4\% \)) and the highest \( S_{1.4 \text{ GHz}} / f_{2\text{ keV}} \) ratio, while the FIRST/RASS-II sample of Brinkmann et al. (2000) is comparable with other soft X-ray selected surveys with similar \( S_{1.4 \text{ GHz}} / f_{2\text{ keV}} \) ratios (e.g. Stocke et al. 1991).

In addition to soft X-ray samples a number of studies have explored the radio properties of hard sources finding much higher X-ray/radio identification rates than those in softer X-ray surveys (Barger et al. 2001; Ciliegi et al. 2003). This is attributed to both observational effects (e.g. deeper radio data) and the insensitivity of both the radio and the hard X-ray wavelengths to obscured AGNs (e.g. Ciliegi et al. 2003). Moreover, hard X-ray selected AGNs with radio counterparts have harder X-ray spectra (i.e. more absorbed) than non-radio detected AGNs (e.g. Bauer et al. 2002; Georgakakis et al. 2004b). It has been suggested that this may be due to the presence of circumnuclear starbursts in radio detected AGNs that both feed and obscure the central black hole (e.g. Bauer et al. 2002). Only two sources in the present sample (\#11, 12) show evidence for absorption with the remaining having soft X-ray spectral properties consistent with little or no photoelectric obscuration. These two sources are also detected in the hard X-ray spectral band (2–8 keV) above the 6\( \sigma \) detection threshold (Georgantopoulos et al. 2004, Paper IV).

Brinkmann et al. (2000) also present follow up spectroscopic observations for about half of their X-ray/radio matched population. They find that \( \approx 71 \pm 5\% \) (324/454) of the sources are AGNs, \( \approx 22 \pm 2\% \) (99/454) are early type galaxies with absorption line spectra, \( \approx 5 \pm 1\% \) (23/454) are classified starbursts and \( \approx 2 \pm 1\% \) (8/454) are clusters. We note that the class of absorption line galaxies in the Brinkmann et al. (2000) sample is likely to be contaminated by AGNs that do not reveal their presence in the optical spectrum (Severgnini et al. 2003).

In our sample about 75\%33\% (9/12) of the X-ray/radio matches are AGNs on the basis of their optical spectra, radio morphology or X-ray/optical properties. Starforming ‘normal’ galaxy candidates represent \( \approx 17 \pm 13\% \) (2/12) of the
X-ray/radio population while we find 1 radio source associated with cluster X-ray emission. We find a larger fraction of star-forming galaxy candidates than Brinkmann et al. (2000) that can be interpreted as tenuous evidence for increasing fraction of ‘normal’ galaxies with decreasing X-ray flux. One should be cautious however, since the estimated fractions above are clearly affected by small number statistics that do not allow firm conclusions to be drawn.

‘Normal’ galaxy candidates have only recently been reliably identified in X-ray surveys suggesting that their number density increases with decreasing X-ray flux (Bauers et al. 2002; Hornschemeier et al. 2003; Kim et al. 2004; Paper II). On the contrary, radio selection below 10 mJy has long been known to be an efficient tool for identifying starburst galaxies albeit with some AGN contamination (Georgakakis et al. 1999; Magliocchetti et al. 2002; Chapman et al. 2003). Star-forming systems appear in increasing numbers with decreasing flux density below 10 mJy and are believed to dominate the sub-mJy radio source counts (Hopkins et al. 1998; 2003). The two ‘normal’ galaxy candidates found in the present study by cross-correlating the X-ray and radio samples have also been identified in Paper II on the basis of both their low $f_X/f_{opt}$ and their optical spectral properties. Therefore, radio selection to the limit of the FIRST survey has identified all the X-ray selected ‘normal’ galaxy candidates in our XMM-Newton/2dF survey. Although the statistics are poor the evidence above suggests that radio selection can potentially single out ‘normal’ star-forming galaxies within X-ray samples.

Magliocchetti et al. (2002) studied the nature of the optically bright ($b_t < 19.45$ mag) FIRST radio sources using optical spectra from the 2dFGRS. They find that as much as 32% of their spectroscopic sample are star-forming systems with the rest being early type galaxies or Seyfert 1 and 2 type systems. The majority of star-forming galaxies are found at relatively low redshifts, $z < 0.1$.

From the 72 FIRST radio sources overlapping with the XMM-Newton/2dF survey (see section 2.2), a total of 18 have optical counterparts on the SDSS brighter than $g \approx 19.2$ mag (i.e. similar to $b_t = 19.45$ mag). On the basis of the Magliocchetti et al. (2002) study from the above 18 radio sources brighter than $g \approx 19.2$ mag we expect about 6 star-forming galaxies. This number can be compared with the two X-ray selected “normal” galaxy candidates identified both in the present study and in Paper II. Although the statistics are poor the analysis above suggests that to the limit of the XMM-Newton/2dF survey we have only identified a sub-sample of the total population of radio selected starforming galaxies (to the limit $g = 19.2$ mag).

It is an interesting exercise to quantify the flux limits that radio and X-ray surveys should have to detect the same population of star-forming galaxies. We use the $L_X - L_{1.4}$ relation of local star-forming galaxies derived by Ranalli et al. (2003) to establish the relation between X-ray flux, $f_X(0.5 - 8$ keV), and radio flux density, $S_{1.4 \text{GHz}}$. Since the Ranalli et al. (2003) $L_X - L_{1.4}$ relation is linear one can factor the distance out to obtain an $f_X - S_{1.4 \text{GHz}}$ relation for star-forming galaxies. This simple transformation however, ignores the k-correction. The effect of the differential X-ray/radio k-correction is nevertheless small compared to the $1\sigma$ rms scatter of the Ranalli et al. (2003) $L_X - L_{1.4}$ relation. The derived $f_X(0.5 - 8$ keV) – $S_{1.4 \text{GHz}}$ relation for star-forming galaxies is compared in Figure 6 with the limits of various X-ray/radio surveys including the XMM-Newton/2dF survey. The region of the parameter space (e.g. X-ray/radio flux range) accessible by these surveys is also marked in this figure. We conclude that our X-ray survey is too shallow to detect the radio selected star-forming population to the limit of the FIRST survey. An order of magnitude improvement in sensitivity at X-ray wavelengths is essential for such a study. Nevertheless, Figure 6 suggests that combining X-ray and radio surveys with carefully selected limiting fluxes is a powerful tool for identifying star-forming galaxies. In particular, this figure shows that the CDF-N has X-ray/radio flux limits that allow a large fraction of the radio/X-ray selected star-forming galaxy population to be identified at both wavelengths. Indeed, Bauer et al. (2002) investigated the association between faint X-ray and radio source populations detected in the 1 Ms CDF-N. They found that the majority of X-ray sources with narrow emission-line spectra also have $\mu$Jy radio counterparts and are likely to be dominated by star-formation activity.

Figure 6. 1.4 GHz radio flux density against X-ray flux in the 0.5–8 keV spectral band. The continuous line is the $f_X - S_{1.4 \text{GHz}}$ relation of local starforming galaxies derived by Ranalli et al. (2003; see text for details). The dashed lines represent the $1\sigma$ rms uncertainty of the $L_X - S_{1.4 \text{GHz}}$ relation. The symbols represent the X-ray flux and radio flux density limits of various X-ray/radio surveys: the open circle is the XMM-Newton/2dF survey with limiting fluxes $f_X(0.5 - 8$ keV) = $10^{-14}$ erg s$^{-1}$ cm$^{-2}$ and $S_{1.4 \text{GHz}} = 1$ mJy, the star is for the Phoenix/XMM survey (Georgakakis et al. 2003b; $f_X(0.5 - 8$ keV) = $10^{-15}$ erg s$^{-1}$ cm$^{-2}$, $S_{1.4 \text{GHz}} = 0.88$ mJy) and the filled circle represents the CDF-N (Bauers et al. 2002; $f_X(0.5 - 8$ keV) = $10^{-16}$ erg s$^{-1}$ cm$^{-2}$, $S_{1.4 \text{GHz}} = 0.05$ mJy). The region of the parameter space (X-ray flux and radio flux density) accessible by these datasets is marked by the dotted diagonal line.
5 CONCLUSIONS

In this paper we explore the nature of the X-ray sources with radio counterparts by cross-correlating the FIRST 1.4 GHz radio survey with a wide field (1.6 deg$^2$) shallow $[f_X(0.5 - 8 \text{ keV}) \approx 10^{-14} \text{ erg s}^{-1}]$ XMM-Newton survey. Our sample comprises 12 X-ray/radio matches representing about 4% of the X-ray selected sample. The majority of these sources (9/12) have properties suggesting AGN activity. Two sources in the sample have narrow emission line optical spectra, X-ray luminosities $L_X \approx 4 \times 10^{41} \text{ erg s}^{-1}$, radio powers $L_{1.4} \approx 5 \times 10^{22} \text{ W Hz}^{-1}$ and X-ray-to–optical flux ratios $\approx -2$ suggesting ‘normal’ galaxies powered by star-formation activity. Finally we find 1 double radio source associated with X-ray cluster emission. Both radio loud and quiet systems in the present sample have mean X-ray spectral properties consistent with $\Gamma \approx 1.9$.

We also argue that radio selection to the limit of the FIRST survey has identified all the X-ray selected ‘normal’ galaxy candidates in our XMM-Newton/2dF survey. Although small number statistics hamper a secure interpretation the evidence above suggests that radio selection can potentially single out ‘normal’ starforming galaxies within X-ray samples. Using the $L_X - L_{1.4} \text{GHz}$ relation for local starforming galaxies we quantify the flux limits that X-ray and radio surveys should have to identify the same population of starbursts.

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APPENDIX A: NOTES ON INDIVIDUAL SOURCES

FIRST J134304.6–000055: this double lobe radio source has optical counterpart lying $\approx 1$ arcsec from the radio centroid. Although optical spectroscopy is not available for this source the radio morphology indicates AGN activity. The X-ray source likely to be associated with FIRST J134304.6–000055 lies 3 arcsec from the radio centroid. The X-ray emission is extended suggesting hot gas cluster emission. Indeed, Couch et al. (1991) used photographic material to identify an optical galaxy overdensity in the vicinity of the X-ray source. Follow-up spectroscopy of selected galaxies in this field by these authors suggests a cluster redshift of $\approx 0.6$. In Paper III we use CCD photometry to identify a statistically significant optical galaxy overdensity in the vicinity of the X-ray source. Using photometric methods (e.g. Postman et al. 1996) these authors estimate a redshift $z \approx 0.6$ for this cluster candidate.

FIRST J134233.0–001553: the radio centroid lies $\approx 3$ arcsec off the X-ray position and $\approx 0.5$ arcsec from the optical counterpart. Optical spectroscopy shows that this source is associated with a broad emission line QSO at a redshift $z = 2.132$.

FIRST J134212.2–001737: the radio position coincides with that of a $g \approx 17.0$ mag galaxy at $z = 0.087$. The X-ray centroid is offset from both the optical and radio source positions by 4 arcsec. This may suggest that the X-ray source may not be associated with either the optical galaxy or the radio source. Keeping this caveat in mind in what follows we assume that both the radio and the X-ray emission originate from the optical galaxy at $z = 0.087$. The 2dF-GRS optical spectrum of this systems suggest an early type spiral with both emission (Hα, [N II]6583Å, [O II]3727Å) and absorption (H+K, Hβ, Hγ, Hδ, NaD5893Å) lines. The X-ray–to–optical flux ratio of $-1.5$ and the X-ray luminosity, $L_X(0.5 - 8) \approx 5 \times 10^{41} \text{ erg s}^{-1}$, suggest either AGN or X-ray properties dominated by stellar processes. In Figure 4 this systems deviates from the mean $L_X - L_{1.4} \text{GHz}$ relation for local star-forming galaxies favoring the AGN scenario for the observed X-ray emission. This source also lies close to the ‘normal’ galaxy region of the $\alpha_{OX} - \alpha_{RO}$ diagram in Figure 5. The C-statistic X-ray spectral analysis suggests soft X-ray spectral properties consistent with Galactic column density absorption. This source is also discussed in Paper II (their source TGN 336Z232). We classify this source LLAGN or ADAF.

FIRST J134347.5+002024: the X-ray and radio positions coincide with 3 arcsec. The optical counterpart is likely to be a $g \approx 19.5$ mag galaxy at a redshift $z = 0.240$. The optical spectrum of this systems does not show emission lines suggesting an early type galaxy. The X-ray–to–optical flux ratio of $+0.3$ and the X-ray luminosity, $L_X(0.5 - 8) \approx 5 \times 10^{41} \text{ erg s}^{-1}$, suggest AGN activity. The two point spectral indices of this source, $\alpha_{OX}$ and $\alpha_{RO}$, are consistent with those of BL-Lacs. This coupled with the soft X-ray spectral properties of this system and the absorption line optical spectrum strongly support the BL-Lac scenario.

FIRST J134414.2+001642: this radio source has double lobe radio morphology indicating AGN activity. The X-ray source likely to be associated with FIRST J134414.2+001642 lies 1 arcsec from the radio cen-
troid. The optical counterpart has \( g \approx 19.7 \) mag and no optical spectroscopic information. The X-ray–to–optical flux ratio of +0.9 is consistent with AGN activity. Our X-ray spectral analysis suggests soft X-ray spectrum. The X-ray–to–optical flux ratio of +0.3 and the X-ray luminosity \( L_X(0.5–8) \approx 8 \times 10^{44} \) erg s\(^{-1}\) are also consistent with AGN activity. This is the only source that has enough photons allowing detailed X-ray spectral analysis. The X-ray spectrum is best fit by a power law spectrum with \( \Gamma \approx 1.8 \) and photoelectric extinction \( N_H < 1.4 \times 10^{20} \) cm\(^{-2}\) consistent with the Galactic value.

**FIRST J134232.4–003151:** the X-ray and radio positions overlap within less than 1 arcsec. The radio source is associated with a \( g = 19.6 \) mag QSO at \( z = 1.209 \). The X-ray–to–optical flux ratio of -0.3 and the X-ray luminosity \( L_X(0.5–8) \approx 8 \times 10^{44} \) erg s\(^{-1}\) are also consistent with AGN activity. This is the only source that has enough photons allowing detailed X-ray spectral analysis. The X-ray spectrum is best fit by a power law spectrum with \( \Gamma \approx 1.8 \) and photoelectric extinction \( N_H < 1.4 \times 10^{20} \) cm\(^{-2}\) consistent with the Galactic value.

**FIRST J134133.4–002432:** the radio source lies 3 arcsec off the X-ray position. The most probable optical counterpart of both the X-ray and radio sources is a \( g \approx 17.1 \) mag galaxy. The X-ray centroid is offset by 3 arcsec from the optical centre of the galaxy. This is a narrow emission line system at \( z = 0.0717 \) showing [O III] 4959 + 5007 Å doublet, Ha and [N II] 6583 Å. Although the signal–to–noise ratio of the spectrum is low, [O II] is not visible in emission. This coupled with the strong [O III] 5007 Å feature may suggest a Seyfert 2 type system. Moreover, the low X-ray luminosity \( L_X(0.5–8) = 1.9 \times 10^{41} \) erg s\(^{-1}\) and X-ray–to–optical flux ratio \( \approx 2.0 \) are consistent with a ‘normal’ galaxy powered by stellar processes. In Figure 4 this system lies close to the mean \( L_X - L_1.4 \) of local star-forming galaxies favoring the star-formation as a probable source of the observed X-ray emission.

**FIRST J134137.7–002555:** The X-ray and radio positions coincide within 3 arcsec. The most probable optical counterpart of both the X-ray and radio sources is a \( g \approx 16.6 \) mag galaxy at \( z = 0.052 \). This narrow emission line galaxy is classified as HII on the basis of its optical spectral properties (Terlevich et al. 1991). Its X-ray luminosity \( L_X(0.5–8) = 1.5 \times 10^{41} \) erg s\(^{-1}\) and X-ray–to–optical flux ratio \( \approx 2.0 \) also suggest ‘normal’ galaxy. Our X-ray spectral analysis also suggests soft X-ray properties. In Figure 4 this system lies close to the mean \( L_X - L_1.4 \) of local star-forming galaxies also favoring the star-formation scenario for the X-ray emission.

**FIRST J134412.9–030006:** the X-ray and radio positions coincide within 3 arcsec. The most probable optical counterpart for both the X-ray and the radio source is a \( g \approx 20.6 \) mag broad line AGN at \( z = 0.708 \). Its X-ray luminosity \( L_X(0.5–8) = 1.4 \times 10^{44} \) erg s\(^{-1}\) and X-ray–to–optical flux ratio \( \approx 0.1 \) also consistent with AGN activity. The C-statistic suggests \( N_H \) consistent with the Galactic column density.

**FIRST J134431.8–002832:** the radio source lies 3 arcsec off the X-ray position. The most probable optical counterpart is a \( g \approx 23.8 \) mag galaxy. No optical spectroscopic information is available for this source. The X-ray–to–optical flux ratio of +0.82 suggests AGN activity. The X-ray spectrum is very steep \( \Gamma \approx 2.3 \) for \( N_H = 2 \times 10^{20} \) cm\(^{-2}\). In the \( \alpha_{OX} - \alpha_{RO} \) diagram this source lies in the borderline between radio loud AGNs and BL-Lacs.

**FIRST J134128.4–003120:** the radio source lies 3 arcsec off the X-ray centroid. The most probable optical counterpart is a \( g = 21.55 \) mag galaxy. Although no optical spectroscopic information is available for this system the X-ray–to–optical flux ratio of +0.65 suggest AGN activity. This source lies in the BL-Lac region of the \( \alpha_{OX} - \alpha_{RO} \) parameter space. The C-statistic suggests hard X-ray spectral properties with photoelectric absorption \( \approx 1.5 \times 10^{22} \) cm\(^{-2}\) \((\Gamma = 1.9)\) which is inconsistent with a BL-Lac.

**FIRST J134447.0–002509:** The X-ray and radio positions coincide within less than 1 arcsec. The radio source has double lobes indicating the presence of an AGN. The optical counterpart has \( g \approx 22.3 \) mag and no optical spectral information are available. The X-ray–to–optical flux ratio of +0.65 also suggest AGN activity. The C-statistic suggests hard X-ray spectral properties with photoelectric absorption \( \approx 10^{22} \) cm\(^{-2}\) \((\Gamma = 1.9)\).

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