Two types of Lyman-α emitters envisaged from hierarchical galaxy formation

この研究は、多層的な銀河形成を想定した場合のリマン-αエミッタの二つのタイプが予想されることが説明されています。

著者名: Shimizu Ikkoh, Umemura Masayuki
誌名: Monthly notices of the Royal Astronomical Society
巻: 406
号: 2
ページ: 913-921
年: 2010-08

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URL: http://hdl.handle.net/2241/107846
doi: 10.1111/j.1365-2966.2010.16758.x
Two types of Lyman $\alpha$ emitters envisaged from hierarchical galaxy formation

Ikkoh Shimizu$^1$ and Masayuki Umemura$^2$

$^1$Institute for the Physics and Mathematics of the Universe (IPMU), The University of Tokyo, 5-1-5 Kashiwanoha, Kashiwa, Chiba 277-8582, Japan
$^2$Center for Computational Sciences, University of Tsukuba, Tsukuba 305-8577, Japan

Accepted 2010 March 30. Received 2010 March 12; in original form 2009 August 4

ABSTRACT

In the last decade, numerous Lyman $\alpha$ (Ly$\alpha$) emitters (LAEs) have been discovered with narrow-band filters at various redshifts. Recently, multiwavelength observations of LAEs have been performed and revealed that while many LAEs appear to be young and less massive, a noticeable fraction of LAEs possess much older populations of stars and larger stellar mass. How these two classes of LAEs are concordant with the hierarchical galaxy formation scenario has not been understood clearly so far. In this paper, we model LAEs by three-dimensional cosmological simulations of dark halo merger in a $\Lambda$ cold dark matter ($\Lambda$CDM) universe. As a result, it is shown that the age of simulated LAEs can spread over a wide range from $2 \times 10^6$ to $9 \times 10^8$ yr. Furthermore, we find that there are two types of LAEs, in one of which the young half-mass age is comparable to the mean age of stellar component, and in the other of which the young half-mass age is appreciably shorter than the mean age. We define the former as Type 1 LAEs and the latter as Type 2 LAEs. A Type 1 LAE corresponds to early starburst in a young galaxy, whereas a Type 2 LAE does to delayed starburst in an evolved galaxy, as a consequence of delayed accretion of a subhalo on to a larger parent halo. Thus, the same halo can experience a Type 2 LAE phase as well as a Type 1 LAE phase in the merger history. Type 1 LAEs are expected to be younger than $1.5 \times 10^8$ yr, less dusty and less massive with stellar mass $M_{\text{star}} \lesssim 5 \times 10^8 M_\odot$, while Type 2 LAEs are older than $1.5 \times 10^8$ yr, even dustier and as massive as $M_{\text{star}} \sim 5 \times 10^8$–$3 \times 10^{10} M_\odot$. The fraction of Type 2s in all LAEs is a function of redshift, which is less than 2 per cent at $z \gtrsim 4.5$, $\sim$30 per cent at redshift $z = 3.1$ and $\sim$70 per cent at $z = 2$. Type 2 LAEs can be discriminated clearly from Type 1s in two-colour diagrams of $z' - H$ versus $J - K$. We find that the brightness distribution of Ly$\alpha$ in Type 2 LAEs is more extended than the main stellar component, in contrast to Type 1 LAEs. This is not only because delayed starbursts tend to occur in the outskirts of a parent galaxy, but also because Ly$\alpha$ photons are effectively absorbed by dust in an evolved galaxy. Hence, the extent of Ly$\alpha$ emission may be an additional measure to distinguish Type 2 LAEs from Type 1 LAEs. The sizes of Type 2 LAEs range from a few tens to a few hundreds kpc. At lower redshifts, the number of more extended, older Type 2 LAEs increases. Furthermore, it is anticipated that the amplitude of angular correlation function for Type 2 LAEs is significantly higher than that for Type 1 LAEs, but comparable to that for Lyman break galaxies (LBGs). This implies that LBGs with strong Ly$\alpha$ line may include Type 2 LAEs.

Key words: galaxies: evolution – galaxies: formation.

1 INTRODUCTION

To explore the early evolutionary phases of galaxies, it is important to understand galaxy formation. Partridge & Peebles (1967) predicted that the starbursts in primeval galaxies emit significant Lyman $\alpha$ (Ly$\alpha$) emission through the recombination of ionized hydrogen in interstellar matter. Although many surveys attempted to discover such Ly$\alpha$ emitting galaxies (Ly$\alpha$ emitters, hereafter LAEs), but did not succeed to find them for a long time. In late 1990s, Cowie & Hu (1998) discovered LAEs with narrow-band filters for the first time. Currently, numerous LAEs have been discovered at high redshifts...
Throughout this paper, we adopt ΛCDM cosmology with the matter density \( \Omega_M = 0.3 \), the cosmological constant \( \Omega_{\Lambda} = 0.7 \), the Hubble constant \( h = 0.7 \) in units of \( H_0 = 100 \text{ km s}^{-1}\text{ Mpc}^{-1} \), the baryon density \( \Omega_B h^2 = 0.02 \) and \( \sigma_8 = 0.92 \) (Spergel et al. 2003).

### 2 MODEL AND NUMERICAL METHOD

#### 2.1 Basic model

To pursue the star formation history in the hierarchical galaxy formation, we simulate the merging history of subgalactic haloes (hereafter subhalo) by three-dimensional cosmological N-body simulations. Here, each particle is regarded as a subhalo that is supposed to consists of dark matter and baryons. We simulate \( N = 256^3 \) subhalos in a comoving volume of \((50 \text{ Mpc})^3\). The mass of a subhalo is \( 2.73 \times 10^8 M_\odot \). It is assumed that the star formation is triggered when a subhalo accretes on to a parent halo. Then we trace the stellar evolution separately for individual subhalos using a spectral synthesis code ‘PEGASE’ (Fioc & Rocca-Volmerange 1997). Moreover, we take into account the effect of dust extinction on Ly\( \alpha \) emission. The present approach allows us to analyse the distributions of star-forming regions in a halo, and also the clustering properties of haloes.

There are basically two types of subhalo accretion. One is the almost contemporaneous accretion of subhaloes in a young small parent halo, and then coeval starbursts take place in the halo. The other is the delayed accretion on to an evolved massive halo, and then a newly triggered starburst and an old stellar population coexist. Both types have the potentiality of becoming LAEs. If they satisfy LAE conditions (see the detail below), we call the former Type 1 LAEs and the latter Type 2 LAEs. A schematic view of Type 1 and Type 2 LAEs is shown in Fig. 1.

#### 2.2 Numerical method

We perform a cosmological N-body simulation with the particle–particle–particle–mesh (P3M) algorithm (Hockney & Eastwood 1981). The numerical scheme is based on Yoshikawa, Jing & Suto (2000). The size of comoving simulation box \( (L_{\text{box}}) \) is set to be the same as the size of LAE survey at \( z = 3.1 \) by Hayashino et al. (2004), that is, \( 50 \text{ Mpc} \) in linear scale. This allows us to adjust LAE conditions by directly comparing with the observation. Here, the periodic boundary condition is imposed. We use the Plummer softening function for gravitational force, with the softening length of \( \epsilon_{\text{grav}} = L_{\text{box}}/(10N^{1/3}) \) \((\sim 20 \text{ kpc} \) in a comoving scale).

A parent halo is found using a friends-of-friends algorithm (Davis et al. 1985) with linking length equal to 0.2 of the mean particle separation. In this study, a system with \( 2.7 \times 10^9 M_\odot \) (say, equal to or more than 100 particles) is identified as a parent halo so that the system mass would be corresponding to observed LAEs. As previously mentioned, each particle (subhalo) in a parent halo has individual age. It is assumed that the star formation is triggered when a subhalo accretes on to a parent halo. The star formation occurs only in the subhalo, and therefore no star formation is triggered at the central of host halo. Furthermore, a subhalo which underwent the star formation once does not trigger the star formation again. Here, the star formation in each subhalo is assumed to occur at the rate as

\[
\psi(t) = f_{\text{int}} \exp \left( -\frac{t}{\tau_\psi} \right). \tag{1}
\]
Two types of LAEs

3 RESULTS

3.1 Two types of simulated LAEs

In Fig. 2, the distributions of mass-weighted age are shown for simulated LAEs. The ages spread widely from $2 \times 10^6$ to $9 \times 10^8$ yr. Interestingly, a significant number of LAEs are much older than $\approx 10^8$ yr, that is, the lifetime of young LAEs predicted by Mori & Umemura (2006) and Shimizu et al. (2007). As seen in Fig. 2, the distributions are fairly continuous from younger LAEs to older ones. Younger LAEs are early coeval starburst galaxies, while older LAEs result from delayed starbursts triggered by later subhalo accretion on to evolved haloes. In order to discriminate delayed starbursts from coeval young starbursts, we plot the young half-mass ages against the mass-weighted mean ages in Fig. 3, where the young half-mass age is defined as the mass-weighted age of the young
half subhaloes included in a host halo. If starbursts are coeval in a halo, young half-mass ages should be comparable to mean ages. It is clearly seen in Fig. 3 that in the part older than \(1.5 \times 10^8\) yr, coeval starbursts disappear and only delayed starbursts appear to be taking place. Hence, in this paper, we define LAEs younger than \(1.5 \times 10^8\) yr as Type 1s, and older ones as Type 2s. It is noted that there is not a clear boundary of two classes at \(1.5 \times 10^8\) yr, but the transition is actually gradual in the sense that coeval and delayed starbursts are blended around \(\approx 10^8\) yr. Nevertheless, as shown below, we find distinctive trends in photometric properties between Type 1s and Type 2s defined here.

Fig. 4 shows EW of \(\text{Ly}\alpha\) emission, \(\text{Ly}\alpha\) luminosity, star formation rate and stellar mass against the mass-weighted age for simulated LAEs. EW decreases with ages for Type 1 LAEs, ranging from 40 to 200 Å, although some are at a level of 350–400 Å. EW for Type 2s is randomly distributed in the range of 30–150 Å. \(\text{Ly}\alpha\) luminosity is basically in proportion to star formation rate for Type 1 LAEs, and gradually decreases with ages. For Type 2 LAEs, \(\text{Ly}\alpha\) luminosity randomly spread in the range of \(\sim 2 \times 10^{42}\) to \(\sim 2 \times 10^{43}\) erg s\(^{-1}\). Interestingly, in old Type 2 LAEs (>6 \(\times 10^8\) yr), a high star formation rate does not always lead to high \(\text{Ly}\alpha\) luminosity. This can be understood by the effect of dust extinction as argued below.

The stellar mass \(M_{\text{star}}\) of LAEs is a fairly monotonic function of age. Type 1 LAEs are less massive with \(M_{\text{star}} \lesssim 5 \times 10^8\) M\(_\odot\), while Type 2s are as massive as \(M_{\text{star}} \sim 5 \times 10^8\)–3 \(\times 10^{10}\) M\(_\odot\). Recent observations show that LAEs detected by rest-frame optical/NIR bands are more massive than \(10^9\) M\(_\odot\) and pretty older (Finkelstein et al. 2007, 2009a; Matsuda et al. 2007; Lai et al. 2008; Uchimoto et al. 2008). Such objects may correspond to Type 2 LAEs.

The fraction of Type 2 LAEs in all LAEs is predicted to be a function of redshift. Fig. 5 represents the Type 2 LAE fraction against redshift. Obviously, the Type 2 fraction increases with decreasing redshift. At \(z \gtrsim 4.5\), the Type 2 fraction is less than 2 per cent, since massive haloes have not grown yet, whereas it is \(\sim 30\) per cent at redshift \(z = 3.1\) and \(\sim 70\) per cent at \(z = 2\). This trend is concordant with the recent observations by Nilsson et al. (2009).

### 3.2 Spectrophotometric properties of simulated LAEs

In Fig. 6, we show SEDs of a typical Type 1 and Type 2 LAE, which are calculated by ‘PEGASE’ (Fioc & Rocca-Volmerange 1997). A Type 1 LAE is composed of young stars, while the host galaxy of a Type 2 LAE is dominated by evolved stars, where a distinct 4000 Å break is seen. Accordingly, the colours of two types of LAEs are different.
3.3 Brightness distribution of a Type 2 LAE

In Fig. 8, the brightness distributions of Lyα emission and those at observed-frame $K$-band flux (which corresponds to rest-frame 6000 Å flux) are shown for a Type 1 LAE at $z = 3.1$. The angular resolution (pixel size) is set to be 1 arcsec. The brightness distributions at observed-frame $K$-band flux trace basically stellar mass distribution. As seen in Fig. 8, both distributions are compact ($\approx 10 \, h^{-1} \, \text{kpc}$), and the extents are quite similar to each other. In Fig. 9, the brightness distributions are shown for a Type 2 LAE at $z = 3.1$. The Lyα emission is diffusely distributed over $\approx 100 \, h^{-1} \, \text{kpc}$ or stronger at the outskirts, while the brightness distributions at observed-frame $K$-band flux exhibit a strong contrast and are more concentrated. Moreover, Type 2 LAE is composed of some clumps. Such clumpy structures can be seen some observed LABs in the SSA22 region (Uchimoto et al. 2008; Webb et al. 2009). This result suggests that some Type 2 LAEs may not be well dynamically

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Type 1 LAE at z=3.1

Figure 8. The brightness distributions of a Type 1 LAE at z = 3.1. Left- and right-hand panels show the brightness distributions of Lyα emission, and those at observed-frame K-band flux (which corresponds to rest-frame 6000 Å flux), of stellar components, respectively. Each colour bar shows the flux of Lyα emission and observed-frame K-band flux, respectively. The angular resolution (pixel size) is set to be 1 arcsec.

Type 2 LAE at z=3.1

Figure 9. Same as Fig. 8, but for a Type 2 LAE at z = 3.1.
Two types of LAEs

Relaxed after mergers, since the duration of strong Lyα emission is shorter than the relaxation time of such a large system.

Since Lyα emission is radiated mostly by starbursts, Lyα bright regions correspond to sites of subhalo accretion. These results suggest that if we observe LAEs in various wavelengths, the extent and morphology of Type 2 LAEs appear to be different in each band. Recently, such an offset has been reported in local starburst galaxy (Hayes et al. 2007; Östlin et al. 2009), while no offset is found in LAEs at a higher redshift of z = 4.4 (Finkelstein et al. 2008). This can be related to the result that the number of more extended, older Type 2 LAEs increases at lower redshifts, as shown below.

Lyα emissions are affected more severely by dust extinction than continuum radiation. The effect of dust extinction is expected to be stronger for old metal-enriched systems. In Fig. 10, the escape fractions of Lyα photons are plotted against ages of LAEs. The escape fractions are a decreasing function of LAE ages as anticipated. In particular, in Type 2 LAEs, the escape fractions decrease down to ≈ 10 per cent. Thus, for aged Type 2 LAEs, Lyα EWs or luminosities become smaller than expected purely from star formation rate (see Fig. 4). Nevertheless, in the outskirts of halo, starbursts by delayed subhalo accretion take place frequently, where the dust extinction is relatively small. Hence, Lyα emissions can be brighter in the outskirts of halo.

As shown in Fig. 9, the size of Type 2 LAEs can exceed 100 kpc in physical size. This is comparable to the size of Lyα blobs (LABs; Matsuda et al. 2004). Thus, some of Type 2 LAEs with EW > 100 Å may account for a part of observed LABs.

3.4 The halo-size distributions of simulated LAEs

Here, we analyse the halo-size distributions of simulated LAEs. Fig. 11 represents the halo-size distributions of all simulated LAEs as a function of mass weighted age at z = 2, 3.1 and 4.5. Here, the radius from the centre of gravity of a halo within which 95 per cent of the total mass is included is defined as the size of a halo. At z = 4.5, there are no Type 2 LAEs. At lower redshifts, Type 2 LAEs appear and the number of more extended, older Type 2 LAEs increases with decreasing redshifts. Interestingly, the range of Type 2 LAE size is quite broad. Small Type 2 LAEs with the size of a few tens kpc are as compact as Type 1 LAEs, while large Type 2 LAEs with the size of a few hundreds kpc are comparable to LABs (Matsuda et al. 2004).

3.5 Clustering properties of two types of LAEs

In order to explore the clustering properties of two types of LAEs, we calculate a two-point angular correlation function (ACF) and a two-point angular cross-correlation function (CCF) between Type 1 and Type 2 LAEs. Fig. 12 represents ACFs of each type LAEs and all simulated LAEs. Furthermore, the ACF of LAEs observed in the SSA22 field (Hayashino et al. 2004) is shown. ACF of Type 1 LAEs as well as ACF of all simulated LAEs is quite weak and well matches the observed ACF in the SSA22 field, whereas ACF of Type 2 LAEs is significantly stronger than that of Type 1 LAEs. This implies that Type 1 LAEs preferentially reside lower density regions at z = 3.1, as shown by Shimizu et al. (2007), while Type 2 LAEs are located in higher density regions. The correlation strength of Type 2 LAEs is comparable to that of LBGs in SSA22 region.
Figure 13. Two-point angular CCF of simulated LAEs. CCF between Type 1 and Type 2 LAEs is depicted by a dash–dot line. For comparison, ACF of all simulated LAEs is also shown by a solid line.

(Giavalisco et al. 1998). This result suggests that Type 2 LAEs may be a subsample of LBGs with strong Ly\textalpha~ emission.

Fig. 13 shows CCF between Type 1 and Type 2 LAEs. The CCF is still weaker than ACF of all simulated LAEs. Very recently, Tamura et al. (2009) have found that the cross-correlation between submm galaxies (SMGs) and LAEs in the SSA22 field is very weak. Type 2 LAEs possess a large amount of dust. If many Type 2 LAEs can be detected as SMGs, our simulation is consistent with this observation.

4 CONCLUSIONS AND DISCUSSION

To explore the origin of two populations of LAEs recently found, we have performed three-dimensional cosmological N-body simulations of subhalo merging history in a ΛCDM universe. We have incorporated star formation history, SED evolution and dust extinction. As a result, we have found that the age of simulated LAEs can spread over a wide range from 2 × 10⁸ to 9 × 10⁸ yr. Furthermore, we have revealed that there are two types of LAEs. We have defined LAEs younger than 1.5 × 10⁸ yr as Type 1s, and older ones as Type 2s. In Type 1 LAEs early coeval starbursts occur in small parent haloes, while in Type 2 LAEs delayed starbursts take place in evolved massive haloes as a consequence of delayed accretion of subhaloes. A parent halo can experience repeatedly a Type 2 LAE phase after a Type 1 LAE phase.

The stellar mass of Type 1 LAEs is \(M_{\text{star}} \lesssim 5 \times 10^8 \, M_\odot\), while Type 2 LAEs are as massive as \(M_{\text{star}} \sim 5 \times 10^8 - 3 \times 10^{10} \, M_\odot\). The physical properties of Type 1 and Type 2 LAEs are concordant with those of two populations of LAEs observed with multiwavelengths (Finkelstein et al. 2007, 2009a; Matsuda et al. 2007; Lai et al. 2008; Uchimoto et al. 2008). The fraction of Type 2s in all LAEs is a function of redshift, which is less than 2 per cent at \(z \gtrsim 4.5\), ~30 per cent at redshift \(z = 3.1\) and ~70 per cent at \(z = 2\). This trend is consistent with two populations of LAEs found by Nilsson et al. (2009). Type 2 LAEs are expected to be discriminated clearly from Type 1 LAEs in two-colour diagrams of \(z' - H\) versus \(J - K\). We find that the brightness distribution of Ly\textalpha~ in Type 2 LAEs is more extended than the main stellar component, in contrast to Type 1 LAEs. This is not only because delayed starbursts tend to occur in the outskirts of a parent galaxy, but also because Ly\textalpha~ photons are effectively absorbed by dust in an evolved galaxy. The sizes of Type 2 LAEs range from a few tens to a few hundreds kpc. At lower redshifts, the number of more extended, older Type 2 LAEs increases. Small Type 2 LAEs are as compact as Type 1 LAEs, while large Type 2 LAEs exceeding 100 kpc are comparable to LABs (Matsuda et al. 2004).

Moreover, we have found that the clustering of Type 2 LAEs is even stronger than Type 1 LAEs. The amplitude of angular correlation function of Type 2 LAEs is comparable to that of LBGs (Giavalisco et al. 1998). This suggests that LBGs with strong Ly\textalpha~ line can be Type 2 LAEs. The two-point angular cross-correlation function is still weaker than that of all LAEs. If many Type 2 LAEs can be detected as SMGs, this result is consistent with recent observation by Tamura et al. (2009).

Interestingly, in a low-redshift universe at 0.2 < \(z < 0.35\), the Galaxy Evolution Explorer (GALEX) have found many LAEs older than \(2 \times 10^8\) yr, more massive than \(10^9 \, M_\odot\) in stellar component, and having small escape fractions of Ly\textalpha~ photons (Deharveng et al. 2008; Finkelstein et al. 2009b). These properties are quite similar to those of Type 2 LAEs we defined in this paper. Hence, local LAEs could be a good sample to study the detailed physical states of Type 2 LAEs.

Furthermore, in an evolved population of LAEs, active galactic nucleus (AGN) events may be anticipated, since supermassive black holes reside evolved galaxies in a local universe (e.g. Marconi & Hunt 2003, and references therein). Recently, Yamada et al. (2009) have found that AGN are exclusively (90 per cent) associated with the massive objects with stellar mass larger than \(10^{10.5} \, M_\odot\). Also, GALEX found that in low-redshift LAEs an AGN fraction is as high as 40 per cent (Finkelstein et al. 2009c). These findings may imply that Type 2 LAEs could have a higher fraction of AGN. Hence, it seems worth investigating the contribution of AGN to Ly\textalpha~ emission especially for Type 2 LAEs (Shimizu & Umemura, in preparation).

ACKNOWLEDGMENTS

We are grateful to T. Hayashino, T. Yamada, Y. Matsuda and R. Yamauchi for providing valuable information and helpful comments, and to M. Mori and R. M. Rich for valuable discussion. Numerical simulations have been performed with computational facilities at Center for Computational Sciences, University of Tsukuba, and the EUP cluster system installed at Graduate School of Frontier Sciences, University of Tokyo. This work was partially supported by the FIRST project based on Grants-in-Aid for Scientific Research (S) (20224002), and also by Grant-in-Aid for Young Scientists (S) (20674003).

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