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Ultraviolet and visible photometry of asteroid (21) Lutetia using the Hubble Space Telescope*

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ABSTRACT

Context. The asteroid (21) Lutetia is the target of a planned close encounter by the Rosetta spacecraft in July 2010. To prepare for that flyby, Lutetia has been extensively observed by a variety of astronomical facilities.

Aims. We used the Hubble Space Telescope (HST) to determine the albedo of Lutetia over a wide wavelength range, extending from ~1500 Å to ~7000 Å.

Methods. Using data from a variety of HST filters and a ground-based visible light spectrum, we employed synthetic photometry techniques to derive absolute fluxes for Lutetia. New results from ground-based measurements of Lutetia’s size and shape were used to convert the absolute fluxes into albedos.

Results. We present our best model for the spectral energy distribution of Lutetia over the wavelength range 1200–8000 Å. There appears to be a steep drop in the albedo (by a factor of ~2) for wavelengths shorter than ~3000 Å. Nevertheless, the far ultraviolet albedo of Lutetia (~10%) is considerably larger than that of typical C-chondrite material (~4%). The geometric albedo at 5500 Å is 16.5 ± 1%.

Conclusions. Lutetia’s reflectivity is not consistent with a metal-dominated surface at infrared or radar wavelengths, and its albedo at all wavelengths (UV-visible-IR-radar) is larger than observed for typical primitive, chondritic material. We derive a relatively high FUV albedo of ~10%, a result that will be tested by observations with the Alice spectrograph during the Rosetta flyby of Lutetia in July 2010.

Key words. minor planets, asteroids: general – minor planets, asteroids: individual: (21) Lutetia

1. Introduction

The Rosetta spacecraft was launched on 2 March 2004 and is heading toward an historic encounter with comet 67P/Churyumov-Gerasimenko (67P/C-G) in 2014. Along the way, Rosetta has flyby encounters with two asteroids: the spacecraft passed at a distance of 803 km from (2867) Steins on 5 September 2008, and an encounter with (21) Lutetia is cur-
Table 1. Log of HST observations of Lutetia.

| Visit Info | Measurements | Objectives |
|------------|--------------|------------|
| ub9h01xxx  | WFPC2/PC, 2 dither points | Near-UV albedo |
| 2008-Nov.-30 | F218W, 4 × 160 s | Bridge from FUV to Visible albedo |
| 17:05-17:45 UT | F255W, 4 × 40 s | |
|            | F360W, 2 × 2 s | |
|            | F606W, 2 × 0.11 s | |
| ub9h02xxx  | ACS/SBC | Far-UV albedo for one hemisphere |
| 2008-Nov.-30 | F140LP, 1 × 1270 s | |
| 18:40-19:25 UT | F165LP, 1 × 1270 s | |
| ub9h03xxx  | WFPC2/PC, 2 dither points | Visible albedo |
| 2008-Nov.-30 | F606W, 2 × 0.11 s | Deep probe for dust debris |
| 20:17-20:58 UT | F606W, 3 × 40 s | Deep probe for companions |
|            | F606W, 7 × 160 s | |
| ub9h04xxx  | WFPC2/PC, 2 dither points | Verify putative companions |
| 2008-Nov.-30 | F606W, 2 × 0.11 s | Repeat of ub9h03xxx |
| 21:52-22:33 UT | F606W, 3 × 40 s | |
|            | F606W, 7 × 160 s | |
| ub9h05xxx  | ACS/SBC | Far-UV albedo for opposite hemisphere |
| 2008-Nov.-30 | F140LP, 1 × 1270 s | |
| 23:27-00:13 UT | F165LP, 1 × 1270 s | |
| ub9h13xxx  | WFPC2/PC, 2 dither points | B-V color of Lutetia and companions |
| 2008-Dec.-15 | F606W, 2 × 0.11 s, 2 × 40 s, 2 × 160 s | |
| 13:26-14:08 UT | F450W, 2 × 0.35 s, 4 × 140 s | |
| ub9h14xxx  | WFPC2/PC, 2 dither points | Verify colors |
| 2008-Dec.-16 | F606W, 2 × 0.11 s, 2 × 40 s, 2 × 160 s | |
| 13:26-14:06 UT | F450W, 2 × 0.35 s, 4 × 140 s | |

In this paper, we focus on the first two objectives, providing the best available pre-flyby estimates of the UV-to-visible spectral energy distribution of Lutetia. The other objectives will be discussed in separate, future publications.

2. Observations, data reduction, and results

We were granted Director’s Discretionary Time (program ID 11957) on HST to perform filter photometry of Lutetia in support of the Rosetta flyby in July 2010. We were initially allocated 5 orbits (synonymous with “visits” in this case) of observing time, which were executed successfully on 30 November 2008, when Lutetia’s heliocentric distance (r) was 2.42 AU, the geometric distance (Δ) was 1.44 AU, and the solar phase angle (φ) was 0.47−0.48. When a preliminary analysis of those data suggested the presence of a previously unknown companion to Lutetia (subsequently determined to be an optical ghost), we were allocated 2 more orbits of HST time, one of which executed successfully on 15 December 2008 (r = 2.45 AU, Δ = 1.50 AU, φ = 7.52) and the other executed successfully one day later (r = 2.45 AU, Δ = 1.50 AU, φ = 7.98). Table 1 provides a log of all the HST observations, detailing the names for the data files, the time span of each visit, the instrument used, the filters employed, the image exposure times, and the objectives of each visit.

For our program, we employed two different instruments: the Planetary Camera (PC) mode of the Wide Field Planetary Camera 2 (WFPC2) was used for the filter photometry covering the near-ultraviolet (NUV) and visible wavelength ranges, while the Solar Blind Channel (SBC) of the Advanced Camera for Surveys (ACS) was used to measure the far-ultraviolet (FUV) flux. The WFPC2/PC has square pixels 0.046 on a side, and the ACS/SBC has (after drizzling to remove geometrical distortion) square pixels 0.025 on a side. In both cases, we used the standard calibrated images from the processing pipeline operated by the Space Telescope Science Institute (STScI).

Because multiple images were obtained for the WFPC2 observations, at two different positions on the detector (to mitigate the effects of cosmic rays and bad CCD pixels), we generally combined them to produce a single, composite image that was used for photometric analysis. Figure 1 shows example composite images for the F218W, F255W, F360W, and F450W filters. The ratio of the signals from the F450W and F606W images taken on Dec. 15 was used to normalize the F450W photometry to the same absolute scale as the photometry taken on Nov. 30. Figure 2 shows all the individual images (not composites) for the short (0.11 s) exposures taken with the F606W filter. The electronic gain for the 0.11 s images taken through F606W filter was half the value used for the other WFPC2 images because Lutetia is bright (V ≈ 10.1), and we needed the extra dynamic range provided by the lower gain setting (0.11 s is the shortest available exposure time for the WFPC2).

The ACS/SBC employs a photon-counting detector that is essentially insensitive to cosmic ray contamination, so we simply took a single, long exposure at a single location on the detector for FUV imaging. The four images (2 for F140LP and 2 for F165LP) are displayed in Fig. 3. Unfortunately, a star passing near Lutetia contaminated the F165LP image taken during visit 05, rendering the photometry unusable for that observation. Although the FUV photometry was obtained ~75 min after the photometry taken through the other filters, we did not attempt to correct the FUV photometry for any light curve variation because that effect was determined to be ~4% (using the same model as Drummund et al. 2010), whereas the statistical error in the F140LP–F165LP photometry is ~17% (see below).

We used standard circular aperture photometry to determine the total signal in each HST filter. We then compared the observed signal to the predicted signal using a model Lutetia
Fig. 1. HST/WFPC2 images of asteroid 21 Lutetia taken through the F218W, F255W, F300W, and F450W filters (see identifying labels on the images). Each image is a 128 × 128 pixel subsection (each image is 5′.89 across) of the full image and is displayed using an asinh intensity stretch (similar to logarithmic) ranging from approximately zero to the maximum intensity in each image. All images have been rotated so that celestial north points up and east is to the left. The scale bar is 1″ across, which subtends 1088 km at Lutetia on December 15 (when the F450W image was taken), and 1044 km on November 30 (when all the other images were taken). Each image is a composite of at least two separate images taken with Lutetia centered at two different locations on the CCD. The “tails” apparent on two of the images are low-level artifacts caused by degraded charge transfer efficiency in the WFPC2 CCD.

albedo spectrum and measured HST throughput curves (a technique called synthetic photometry). We iteratively adjusted the albedo model until the measured count rates from the HST filter photometry matched the predicted values from synthetic photometry to within the measurement uncertainties. In addition to the Hubble photometry, we used a ground-based spectrum taken on 28 November 2008 (Perna et al. 2010), just two days prior to the first HST observations, to constrain the slope of Lutetia’s albedo at visible wavelengths. All albedos discussed here refer to the geometric albedo, which is the albedo at a solar phase angle of 0°. We adopted a phase correction factor of 0.91 when converting fluxes from a phase angle of 0° to the observed phase angle of 0°.48 on 2008 November 30, which is consistent with the phase law deduced from visible observations (Belskaya et al. 2010).

In coordination with our HST effort, we acquired high-angular resolution adaptive-optics imaging at the Keck Telescope near the times of the 2008 opposition. We combined these data with data previously acquired at the Very Large Telescope (VLT) near the prior opposition, in June 2007, to provide improved estimates of the size, shape, and spin axis of Lutetia, and also to search for satellites (Carry et al. 2010; Drummond et al. 2010; Merline, in prep.). We adopt here the hybrid shape model derived by combining the results from the direct size measurements and the inversion of various light curve data (Carry et al. 2010; Drummond et al. 2010). This model
is then used to predict Lutetia’s projected area at the time of the HST observations. The predicted projected area varies between 9615 km$^2$ and 10081 km$^2$ over the course of the HST observations, and this variation of $\sim$4.8% is within the expected uncertainty in the prediction itself ($\sim$6%). We have adopted a projected area of 10,000 km$^2$, corresponding to an effective diameter for Lutetia of 113 km, to set the absolute scale for the albedo. (Note that the mean 3-dimensional diameter of Lutetia is 105 km, as determined by Drummond et al. 2010. But when viewed from high latitudes, or close to pole-on, the effective diameter of the projected disk is larger, due to viewing primarily the a,b dimensions. An independent size estimate can be derived from the HST images because Lutetia is slightly resolved, but detailed modeling is required to derive accurate size and shape estimates, and we are deferring this work to a future paper.) We estimate that the uncertainty in the albedo derived below is $\sim$6%.

We used the solar spectrum from Colina et al. (1996) to convert from albedo to absolute fluxes. Table 2 summarizes the numerical results after the final iteration, and Fig. 4 displays our best estimate for Lutetia’s absolute flux at the time of the first HST observations on 30.726 November 2008.

Applying synthetic photometry to the HST data is essential for obtaining accurate results because Lutetia’s flux varies dramatically across the passbands of all the filters used, except $F450W$ and $F606W$. Figure 5 shows the predicted count rates as a function of wavelength for each of the filters employed during the HST observations, as well as the response for the difference $F140LP$–$F165LP$. Determination of the FUV flux is especially
Fig. 3. HST/ACS images of asteroid 21 Lutetia taken through the F140LP and F165LP filters (see identifying labels). Each image is displayed using an asinh intensity stretch (similar to logarithmic) ranging approximately from zero to the maximum intensity in each image. All images have been rotated so that celestial north points up and east is to the left. The scale bar is 0.5' across, which subtends ~520 km on 30 November 2009 when the images were taken (start times are labeled on each image).

Table 2. HST photometry of Lutetia.

| Filter     | Measured signal (e s$^{-1}$) | Model signal (Total e s$^{-1}$) |
|------------|------------------------------|--------------------------------|
|            | Total                        | Error                          |          |
| F140LP     | 12.52                        | 0.10                           | 12.52    |
| F165LP     | 11.69                        | 0.10                           | 11.64    |
| F140LP-F165LP | 0.828                      | 0.141                          | 0.879    |
| F218W      | 33.7                         | 1.7                            | 33.4     |
| F255W      | 226                          | 7.0                            | 226      |
| F300W      | 9625                         | 150                            | 9499     |
| F450W      | 120,160                      | 2400                           | 122,990  |
| F606W      | $1.107 \times 10^6$         | $4.42 \times 10^4$            | $1.122 \times 10^6$ |

problematic owing to “red leak” issues (i.e., when much, or even most, of the observed signal is produced by photons whose wavelengths are much longer than the wavelength of the peak in the filter transmittance), which Fig. 5 graphically illustrates. For the F140LP filter, only ~10% of the observed signal is produced by photons having wavelengths shortward of 1895 Å. For F165LP, only ~10% of the observed signal is produced by photons having wavelengths shortward of 1975 Å. The 50% point (i.e., half the observed signal is produced by photons either shortward or longward of that wavelength) occurs at 3340 Å for F140LP and 3390 Å for F160LP. We mitigate the red leak problem by forming the difference signal (F140LP−F165LP), which significantly improves the situation, as indicated in Fig. 6. In that case, ~40% of the difference signal is produced by photons having wavelengths shortward of 1675 Å, and the 50% point occurs at 2400 Å. Nevertheless, the red leak remains a major issue affecting the accuracy of our FUV results.
Fig. 4. Our best estimate for the Lutetia spectrum at the time of the HST observations on 30 November 2008.

In order to match the measured signals for the $F140LP$ and the $F165LP$ filters, and additionally match their difference signal ($F140LP-F165LP$), we had to increase the system throughput (QT) for each of those filters by a factor of 2.5 for wavelengths longward of 2000 Å, relative to the curves currently adopted by the STScI (i.e., the red leak is even worse than originally thought). After consultation with the relevant experts at the STScI, we concluded that the large uncertainty in the long wavelength response of the FUV filters justifies our empirical approach. While our choice for the filter throughputs is certainly not unique, and may be incorrect in detail, we are confident in our assessment that a red leak adjustment for the $F140LP$ and $F165LP$ filters, of approximately the magnitude adopted here, is required to produce consistent results.

Additional information on Lutetia’s UV albedo is available from IUE observations performed in 1982, originally analyzed by Roettger & Buratti (1994). We obtained these data from the STScI archive, and reanalyzed them using new information on the phase behavior (Belskaya et al. 2010) and size of Lutetia (Drummond et al. 2010). The IUE spectrum (Fig. 7) is noisy and was taken at a solar phase angle of 26°, making the correction to geometric albedo rather uncertain. Adopting a phase correction factor of 3.1, which is the value observed for Lutetia at visible wavelengths, we obtain an average albedo of ~0.14 near 2670 Å. The latter is double the value adopted by Roettger & Buratti (1994) (0.074, after correction to the new effective diameter), apparently owing to their use of a different phase law. The phase law typically depends on both the absolute albedo and the wavelength. Thus, it is not surprising that our phase correction factor is significantly different than the one adopted by Roettger & Buratti (1994), especially since there is scant data available on the phase behavior of Lutetia at UV wavelengths. In order to match the HST photometry from the $F255W$ filter, we adopted an albedo of ~0.10 near 2670 Å, which is approximately halfway between the two different IUE results.

Of greater concern is the slope of Lutetia’s albedo between 2400 Å and 3300 Å. The HST data suggest there is a sharp drop in Lutetia’s albedo in the wavelength range ~3000–3300 Å, while the IUE data indicate that the albedo is essentially constant over the wavelength range ~2400–3300 Å. Although the IUE observations were performed nearly 27 years before the HST observations, we would not expect the UV albedo of Lutetia to change either in its slope or its absolute value over that time. The aspect angles of the HST and IUE observations were significantly different (Lutetia’s sub-Earth latitude was ~73° during the HST observations and ~51° during the IUE observations), and perhaps albedo variation over Lutetia’s albedo could explain the differences in the HST and IUE results. However, the long exposure time for the IUE spectrum (3.2 h) covered nearly 75% of Lutetia’s lightcurve period, suggesting that surface variation may not play a significant role. In summary, there appears to be a discrepancy between the IUE and HST results for the slope of Lutetia’s UV albedo near 3100 Å. Nevertheless, both the IUE and HST data indicate that the NUV albedo is significantly smaller than the visible albedo. The HST data further suggest that the FUV albedo is approximately 60% of the visible albedo.

3. Discussion

Lutetia was extensively observed in the 1970s, yielding visible and near-IR reflectance spectra (McCord & Chapman 1975), radiometric albedos and diameter estimates (Morrison 1977), and polarimetric albedos and diameter estimates (Zellner & Gradie 1976), which have been confirmed by similar observations reported during the last decade (see Belskaya et al. 2010). Based
Fig. 5. After adopting our best estimate for Lutetia’s flux, we plot the predicted count rates as a function of wavelength for each of the filters employed during the HST observations. “F140LP–F165LP” refers to the difference between the F140LP and F165LP filters. For clarity, the F140LP curve is not explicitly labeled, but it is essentially identical to the F165LP curve longward of 1650 Å and essentially identical to the F140LP–F165LP curve shortward of that wavelength. Note the logarithmic scale.

Fig. 6. The cumulative fractional contribution to the observed count rate as a function of wavelength for the difference (F140LP–F165LP) of the two far-UV filters employed during the HST observations. Approximately 40% of the observed signal is produced by photons having wavelengths smaller than 1675 Å, but 50% of the signal is coming from light longward of 2400 Å.
Fig. 7. Albedo spectrum of Lutetia derived from IUE observations made on 7 January 1982 at a solar phase angle of 26.1. The albedo at 2670 Å adopted by Roettger & Buratti (1994) is shown by the dashed green line (the lower dashed line), while our equivalent value, using a different phase correction, is depicted by the dashed red line (the upper dashed line). The blue curve is a Fourier-filtered version of the IUE spectrum, passing only the lowest 1% of spatial frequencies.

Numerous researchers in the last few years (Barucci et al. 2005, 2008; Lazzarin et al. 2009; Perna et al. 2010; see summary by Belskaya et al. 2010) have argued that Lutetia shows certain spectral characteristics (e.g., in the thermal IR) that resemble several CO and CV meteorites, but not an iron meteorite. However, mineralogical interpretations of thermal IR spectra must be made cautiously because particle size, in addition to composition, can strongly affect the observed spectral features (Vernazza et al. 2010). We further note that: (a) the lack of a drop-off in Lutetia’s spectral reflectance below 0.55 μm and its relatively high albedo make it inconsistent with CV meteorites (see Gaffey 1976, for instance); and (b) CO meteorites display a 1 μm olivine band that is absent in Lutetia’s reflectance spectrum (see Fig. 3 of Barucci et al. 2005).

It was first suggested by Chapman & Salisbury (1973) that what we now term an M-type spectrum might be associated with enstatite chondrites (ECs). More recently, Rivkin et al. (2000) have suggested that a hydrated EC is a plausible composition for Lutetia, consistent with the recent analysis of Vernazza et al. (2009). From rotationally resolved visible and near-IR spectra of Lutetia, Nedelcu et al. (2007) claimed a better match with CC in one hemisphere and with EC in the other, but this hemispherical spectral asymmetry has not yet been confirmed by other researchers.

Recent dynamical work (Baer et al. 2008; Fienga et al. 2009) has provided an estimate for Lutetia’s mass, which combined with the new size estimates (Drummond et al. 2010; Carry et al. 2010) yield a bulk density of ~4 g cm⁻³ (the formal uncertainty ranges from 2.4–5.1 g cm⁻³; see Drummond et al. 2010). This
density is too small for an object having a dominantly metal component and seems more compatible with an EC-like composition (Drummond et al. 2010).

As discussed by Roettger & Buratti (1994), the slope of Lutetia’s NUV albedo is similar to that of the M- and S-type asteroids observed by IUE. The albedo of the C-type asteroids observed by IUE increases (by ~10%) between 2400 Å and 3000 Å, whereas there is little to no variation in Lutetia’s albedo over this wavelength range. Both the IUE and HST data demonstrate that the absolute value of Lutetia’s NUV albedo is larger than typically observed for the C-type asteroids.

According to the HST data, Lutetia has a rather high FUV albedo of ~10% over the entire wavelength range from ~1500 Å to ~3000 Å. This can be compared to an FUV albedo of ~4% for the Earth’s Moon (Henry et al. 1995) and the E-type asteroid 2867 Steins (A’Hearn et al. 2010), the only asteroid observed at FUV wavelengths. Furthermore, neither the Moon nor Steins show an abrupt drop in albedo near 3200 Å.

We compared Lutetia’s albedo to laboratory reflectivity measurements of a wide variety of materials, including meteorite and lunar samples (Wagner et al. 1987), and none of them appear to be good analogs for Lutetia’s surface. The feldspar powders have a sharp albedo drop near 3000 Å and have NUV-FUV albedos similar to that of Lutetia, but the ratios between their visible and UV albedos are several times larger than we find for Lutetia. SO2 frost also has a sharp drop in albedo near 3000 Å and an FUV albedo in the range of 10−15%, both of which are consistent with Lutetia’s UV spectrum. However, the visible-to-UV albedo ratio of SO2 frost is several times larger than Lutetia’s, and exposed frost isn’t expected to be present on Lutetia’s surface. The albedos of chondritic meteorite samples are significantly smaller than Lutetia’s albedo at all wavelengths, in addition to not matching Lutetia’s spectral variations. Similarly, lunar samples tend to have lower UV albedos than Lutetia. Spectra of various mineral powders (e.g., iron, clays, sulfur) also show striking differences when compared to Lutetia’s spectrum. Perhaps some mixture of samples could be found to approximate Lutetia’s spectrum, but such an effort is beyond the scope of this paper. We note, however, that many of the materials measured by Wagner et al. (1987) have spectral features shortward of 2000 Å, which are potentially observable by the Alice instrument during the Rosetta flyby.

4. Conclusion

Using the Hubble Space Telescope, we measured the albedo of asteroid (21) Lutetia over a wide wavelength range, extending from the far ultraviolet (~1500 Å) to the visible (~7000 Å). The HST results reported here suggest a sharp drop in Lutetia’s albedo near 3100 Å, and an essentially constant FUV albedo of ~10% between 1400−3000 Å. The absolute value and spectral variation of Lutetia’s UV-visible albedo is not well-matched by the spectra of any meteorites measured in the laboratory. Lutetia may well be composed of material that is either rare or not yet represented in our meteorite collections.

Fig. 8. Our best estimate for the albedo of Lutetia is plotted as a solid black line. The spectrum plotted in red is from a ground-based observation taken on 28 November 2008. The blue rectangle shows the wavelength coverage and the range of possible albedos derived from an IUE spectrum taken on 7 January 1982. The HST data (×) are plotted at the wavelengths where the predicted count rate is largest; the horizontal bars give the “effective bandwidth” of the filters, as specified in the STScI Instrument Handbooks. The wavelengths of the standard UBVR bands are also displayed for reference. Only the error bar for the F140LP–F165LP difference filter case is displayed because the error bars for the other filters are smaller than the plotting symbols.
Lutetia’s FUV albedo is considerably higher than the values measured for C-chondrites and the Earth’s Moon (∼4%), which implies that Lutetia should be a relatively easy target for the Alice instrument when it makes observations during the Rosetta close flyby in July 2010.

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