The REBELS ALMA Survey: cosmic dust temperature evolution out to $z \sim 7$

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ABSTRACT

ALMA observations have revealed the presence of dust in the first generations of galaxies in the Universe. However, the dust temperature $T_d$ remains mostly unconstrained due to the few available FIR continuum data at redshift $z > 5$. This introduces large uncertainties in several properties of high-$z$ galaxies, namely their dust masses, infrared luminosities, and obscured fraction of star formation. Using a new method based on simultaneous [C ii] $158\mu$m line and underlying dust continuum measurements, we derive $T_d$ in the continuum and [C ii] detected $z \approx 7$ galaxies in the ALMA Large Project REBELS sample. We find $39 \text{K} < T_d < 58 \text{K}$, and dust masses in the narrow range $M_d = (0.9 - 3.6) \times 10^7 M_\odot$. These results allow us to extend for the first time the reported $T_d(z)$ relation into the Epoch of Reionization. We produce a new physical model that explains the increasing $T_d(z)$ trend with the decrease of gas depletion time, $t_{\text{dep}} = M_\text{d} / S\text{FR}$, induced by the higher cosmological accretion rate at early times; this hypothesis yields $T_d \propto (1 + z)^{0.4}$. The model also explains the observed $T_d$ scatter at a fixed redshift. We find that dust is warmer in obscured sources, as a larger obscuration results in more efficient dust heating. For UV-transparent (obsured) galaxies, $T_d$ only depends on the gas column density (metallicity), $T_d \propto N_H^{1/6}$ ($T_d \propto Z^{-1/6}$). REBELS galaxies are on average relatively transparent, with effective gas column densities around $N_H = (0.03 - 1) \times 10^{21} \text{cm}^{-2}$. We predict that other high-$z$ galaxies (e.g. MACS0416-Y1, A2744-YD4), with estimated $T_d \gg 60 \text{K}$, are significantly obscured, low-metallicity systems. In fact $T_d$ is higher in metal-poor systems due to their smaller dust content, which for fixed $L_{\text{IR}}$ results in warmer temperatures.

Key words: galaxies: high-redshift, infrared: ISM, ISM: dust, extinction, methods: analytical – data analysis

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1 INTRODUCTION

The rest-frame Ultraviolet (UV) emission from galaxies in the Epoch of Reionization (EoR) has been extensively studied thanks to the Hubble Space Telescope (HST) and ground-based telescopes (for UV luminosity functions see: Bradley et al. 2012; Oesch et al. 2013; McLure et al. 2013; Bowler et al. 2015; Atek et al. 2015; Bouwens et al. 2015; Livermore et al. 2017, for single detections see: Bouwens et al. 2010, 2011; Ellis et al. 2012; Bradley et al. 2014; Oesch et al. 2016).

Recently, also the Far-Infrared (FIR) emission from such early sources has become observable thanks to the advent of high sensitivity millimetre interferometers such as the Atacama Large Millimeter Array (for surveys see: Capak et al. 2015; Carilli et al. 2016; Bouwens et al. 2016; Barisic et al. 2017; Bowler et al. 2018; Bethermin et al. 2020; Schaerer et al. 2020, for comprehensive reviews see: Carilli & Walter 2013; Dunlop 2016). The combination of data from these different instruments has allowed us to glimpse the interstellar medium (ISM) of early galaxies (see e.g. Dunlop et al. 2013; Stark 2016; Dayal & Ferrara 2018; Hodge & da Cunha 2020) for the first time.

The FIR emission from galaxies includes FIR lines, arising from atomic and molecular species in the ISM, and dust continuum emission. One of the brightest (and thus most commonly observed) FIR lines is the fine-structure transition of singly ionized carbon $[\text{C} \text{\textsc{i}}] 158\mu\text{m}$, which traces mainly the neutral atomic gas in the ISM (Stacey et al. 1991; Hollenbach & Tielens 1999; Wolfire et al. 2003). The dust continuum is the thermal radiation emitted by dust grains heated by the UV and optical light coming from young stars (see e.g. Draine 1989; Meurer et al. 1999; Calzetti et al. 2000; Weingartner & Draine 2001; Draine 2003).

Dust grains span a wide range of physical temperatures deepening on their own physical properties and the radiation field heating them (see e.g. Draine 2003). Nevertheless, under the approximation of thermal equilibrium, the FIR spectral energy distribution (SED) can be well approximated by an isothermal grey-body function (Hildebrand 1983). Where available, observations at rest-frame Mid-Infrared (MIR) wavelengths, short-wards of the FIR emission peak, have shown significant deviations from a grey-body (Dunne & Eales 2001; Blain et al. 2003; Kovács et al. 2010; Casey 2012; Dale et al. 2012; Galametz et al. 2012; Kirkpatrick et al. 2012; da Cunha et al. 2008; da Cunha et al. 2015; Casey et al. 2018; Reuter et al. 2020). Unfortunately, MIR wavelengths are inaccessible at $z > 5$ with current instruments. Given the limited availability of observations sampling the full dust emission regime, the most common and physically-sound approach is to assume isothermal dust emitting as a grey-body. The key dust properties constrained through SED fitting are the dust temperature $T_d$ and mass $M_d$, which are degenerate quantities.

Reliably determining high-$z$ galaxies $T_d$ and $M_d$ holds the key to many problems related to early galaxy formation and evolution. For example, it might shed light on the heating produced by obscured star formation, and on the nature and processes governing the dust formation and content in the early universe.

Interestingly, in the last few years several works spanning the redshift range $0 \leq z \leq 6$ have suggested the presence of a direct correlation between $T_d$ and redshift (see e.g. Magdis et al. 2012a; Magnelli et al. 2013; Béthermin et al. 2015; Schreiber et al. 2018; Faisst et al. 2020; Bouwens et al. 2020; Reuter et al. 2020). Combining the few available dust temperature estimates for individual galaxies at $z > 5$ with lower-$z$ stacking results, different works came to discrepant conclusions. While Bouwens et al. (2020) have confirmed the reported linear $T_d(z)$ increase, Faisst et al. (2020) suggested instead a flattening of the relation at higher redshift. A physical and quantitative interpretation of either trend is still lacking. Previous works suggested that warm dust temperatures at high-$z$ might result from a more compact dust geometry in high-$z$ galaxies (w.r.t. local galaxies), with dust being mostly located in the vicinity of young stars, and thus being more efficiently heated (Liang et al. 2019; Sommovigo et al. 2020).

Other works have suggested that the increasing $T_d(z)$ trend can be qualitatively ascribed to the growing specific star formation rate at high-$z$ (Magnelli et al. 2014; Ma et al. 2016, 2019; Shen et al. 2021), together with the lower dust content of early galaxies (see e.g. Aoyama et al. 2017; Behrens et al. 2018; Shen et al. 2021; Pallottini et al. 2022). A lower dust content would in fact result in warmer dust temperatures for a given FIR luminosity $L_{\text{FIR}}$.

 Firmer conclusions on the $T_d$ evolution at early times have been hindered by two factors. First, the number of galaxies at $z > 5$ for which multiple FIR continuum observations are available is very limited (Hashimoto et al. 2019; Behrens et al. 2018; Harikane et al. 2020; Bakx et al. 2020; Faisst et al. 2020). In fact, most ALMA sources, when detected in dust continuum, have only a single (or very few) data point(s) at FIR wavelengths (e.g. Knudsen et al. 2016a; Bouwens et al. 2016; Pavesi et al. 2016a; Barisic et al. 2017; Bowler et al. 2018; Pavesi et al. 2019; Hashimoto et al. 2019; Tamura et al. 2019; Bethermin et al. 2020). Hence a value for $T_d$ is often assumed a priori in the fitting procedure (or associated to very large experimental uncertainties). The second limiting factor is the very large scatter of measured $T_d$ values, ranging from $T_d \simeq 25$ K (Harikane et al. 2020) up to $T_d > 80$ K in the narrow redshift range $z = 6.2 - 8.3$ (Bakx et al. 2020).

The lack of (or poor) knowledge of $T_d$ at high-$z$ results in very large uncertainties on $M_d$ as well as derived galaxy properties, such as infrared luminosity ($L_{\text{IR}}$) and obscured star formation rate (SFR; see e.g. Sommovigo et al. 2020). In the future, the problem can be mitigated thanks to further ALMA observations in multiple higher-frequency bands (Bands 7,8,9). In fact, for galaxies at $z \lesssim 5$ such observations sample the SED closer to the emission peak in the FIR (see Bakx et al. 2021).

To overcome the current FIR observational limitations at $z \gtrsim 5$, in Sommovigo et al. (2021) we developed a new method aimed at simultaneously constraining $T_d$ and $M_d$ with a single band measurement. The main idea is to combine the widely observed fine-structure $[\text{C} \text{\textsc{i}}] 158\mu\text{m}$ line with the underlying dust continuum emission at the same frequency (1900 GHz).

Our method can improve the reliability of the interpretation of $[\text{C} \text{\textsc{i}}]$ and continuum observations from millimetre interferometers. This is particularly relevant in the context of recent ALMA large programs targeting $[\text{C} \text{\textsc{i}}]$ emitters at $z \gtrsim 5$, such as the very recent "Reionization Era Bright Emission Line Survey" (REBELS; PI: Bouwens, Bouwens et al. 2021). REBELS studied 40 of the brightest known galaxies

An additional parameter is the dust emissivity, $\beta_d$; see Section 2.
at $z > 6.5$ identified over a 7 deg$^2$ area of the sky, systematically scanning for bright ISM-cooling lines, [C II] at 65.7 – 7.7 (−21.3 < $M_{\text{CII}} < −22.5$); they are luminous galaxies with stellar masses$^2$ in the range $10^9 < M_*/M_\odot < 10^{10}$, and relatively high SFRs around $20 < \text{SFR}/M_\odot\text{yr}^{-1} < 200$ (Bouwens et al. 2021).

REBELS galaxies are UV-selected sources at redshift $z = 6.5 – 7.7$ (−21.3 < $M_{\text{CII}} < −22.5$); they are luminous galaxies with stellar masses$^2$ in the range $10^9 < M_*/M_\odot < 10^{10}$, and relatively high SFRs around $20 < \text{SFR}/M_\odot\text{yr}^{-1} < 200$ (Bouwens et al. 2021).

Here we apply our method (Sommovigo et al. 2021) to the 13 REBELS targets detected both in continuum (3$\sigma$) and in [C II] (5$\sigma$) at $z = 6.5 – 8.5$ (for the sources properties see Table 1). They represent the first statistical sample of continuum detections at such early epochs, featuring a four-fold increase of the size of the previously available galaxy sample.

Constraining the dust temperatures and masses of REBELS galaxies is crucial for a better understanding of SF obscuration and dust production in these early galaxies. Moreover, our analysis on the REBELS sample can investigate for the first time the reported cosmic dust temperature evolution (Magdis et al. 2012a; Magnelli et al. 2013; Béthermin et al. 2015; Schreiber et al. 2018; Faisst et al. 2020; Bouwens et al. 2020) into the Epoch of Reionization. In particular, the eventual flattening in the $T_d(z)$ evolution would be more robustly identified at the high ($z = 7$) redshift of the REBELS sample.

This paper is organised as follows. In Section 2 we summarize the Sommovigo et al. (2021) method to compute $T_d$ which is then applied to REBELS galaxies in Section 3. Section 4 is devoted to the analysis of the reported $T_d$-redshift relation in the light of the derived REBELS results. There we introduce a new physical model to explain the observed increasing $T_d$–$z$ trend. In Section 5 a summary and discussion of the results is given.

Throughout the paper, we assume a ΛCDM model with the following cosmological parameters: $\Omega_M = 0.3075$, $\Omega_\Lambda = 1 – \Omega_M$, $\Omega_b = 0.0486$, $\sigma_8 = 0.81$. $\Omega_M$, $\Omega_\Lambda$, $\Omega_b$ are the total matter, vacuum, and baryonic densities, in units of the critical density; $h$ is the Hubble constant in units of 100 km s$^{-1}$ Mpc$^{-1}$, and $\sigma_8$ is the late-time fluctuation amplitude parameter (Planck Collaboration et al. 2016).

2 METHOD

In Sommovigo et al. (2021) we proposed a novel method to derive the dust temperature in galaxies, based on the combination of the 1900 GHz continuum and the super-imposed [C II] line emission. We summarize the method in the following.

We use the [C II] luminosity, $L_{\text{CII}}$, as a proxy for the total gas mass $M_g$, or equivalently the dust mass $M_d$, given a dust-to-gas ratio $D$:

$$M_d = DM_g = D\alpha_{\text{CII}}L_{\text{CII}}.$$  

(1)

where $\alpha_{\text{CII}}$ is the [C II]-to-total gas conversion factor.

We derive an analytic expression for $\alpha_{\text{CII}}$ using empirical relations such as the Kennicutt–Schmidt relation (Kennicutt 1998, hereafter, KS), and the De Loore relation between $L_{\text{CII}}$–SFR (De Loore et al. 2014, hereafter, DL). This yields

$$\alpha_{\text{CII}} = 32.47 \frac{y^2}{k^\beta_{\text{CII}}^\mu} \frac{M_\odot}{T_\odot},$$  

(2)

where $\Sigma_{\text{SFR}}$ is the SFR surface density, and $k$ is the “burstiness parameter” (Ferrara et al. 2019; Pallottini et al. 2019; Vallini et al. 2020) which quantifies upwards deviations from the KS relation ($k > 1$ for starbursts$^3$). The factor $y = r_{\text{CII}}/r_*$ is introduced since there is growing evidence (1.5 < $y < 3$ at $z > 4$, see e.g. Carniani et al. 2017, 2018; Matthee et al. 2017, 2019; Fujimoto et al. 2019; Rybak et al. 2019; Fujimoto et al. 2020; Ginolfi et al. 2020; Carniani et al. 2020) that $z > 4$ [C II] emission (size $r_{\text{CII}}$) is more extended than the UV size ($r_*$).

We assume the dust-to-gas ratio $D$ to scale linearly with metallicity down to $Z \lesssim 0.1 Z_\odot$ (James et al. 2002; Draine & Li 2007; Galliano et al. 2008; Leroy et al. 2011; Rémy-Ruyer et al. 2014):

$$D = D_0 \left( \frac{Z}{Z_\odot} \right),$$  

(3)

where $D_0 = 1/162$ is the Galactic value (Rémy-Ruyer et al. 2014). This simple linear $D – Z$ relation is confirmed to hold in simulated galaxies at $z \sim 7$ for stellar masses in the range $10^9 < M_*/M_\odot < 10^{11}$ (Dayal et al., in prep. 2022, see also Ma et al. 2016; Torrey et al. 2019).

Armed with an expression for $\alpha_{\text{CII}}$, $D$, and thus $M_d$ (eq. 1), we can constrain $T_d$ using the continuum flux at 1900 GHz. We consider Milky Way-like dust$^4$, for which standard values for the dust opacity$^5$ are $\kappa_\nu = \kappa_{\text{CII}} (\nu/\nu_{\text{CII}})^{\beta_{\text{CII}}} = (10.41 \text{ cm}^2 \text{g}^{-1}, 1900 \text{ GHz}, 2.03)$ (Weingartner & Draine 2001; Draine 2003).

In Sommovigo et al. (2021) we rewrote the equation for the continuum flux emitted by a given dust mass$^6$ with (CMB-corrected) equilibrium temperature $T_d$, yielding the following explicit expression for $T_d$:

$$T_d = \frac{T_0}{\ln(1 + f^{-1})},$$  

(4)

where $T_0 = h_P
v_0/k_B = 91.86$ K is the temperature corresponding to the [C II] transition energy at $v_0 = 1900$ GHz, $k_B$ and $h_P$.

$^2$ derived from rest-frame UV SED fitting by Stefanon et al., in prep. (2022) using BEAGLE (Chevallard & Charlot 2016). The authors adopt a constant Star Formation History (SFH), 0.2 $Z_\odot$ metallicity, a Calzetti et al. (2000) dust extinction law, and a Chabrier (2003) 0.1 – 300 $M_\odot$ initial mass function. Note that the correction on the given $M_*$ values, required to be consistent with a Salpeter 1 – 100 $M_\odot$ IMF (used in the rest of the paper), is well within their uncertainties and does not affect significantly our results. We caution that using a non-parametric prescription for the SFH might result in $M_*$ values up to a factor $\sim$3 larger (for a detailed discussion on this point see Topping et al., in prep. 2022). Implications on the results presented here are discussed in the text.

$^3$ $\kappa_\nu = 1$ for normal galaxies, being defined as $\kappa_\nu = \Sigma_{\text{SFR}}/(10^{-12} \text{ cm}^2 \text{g}^{-1})$ (Heiderman et al. 2010)

$^4$ Milky Way-like dust seems the best suited to reproduce high-z sources properties, see (Bowler et al. 2018; Schouws et al. 2021; Ferrara et al., in prep. 2022) for a detailed discussion.

$^5$ This power-law approximation for the dust opacity is valid at wavelengths $\lambda > 20 \mu$m, well within the FIR range.

$^6$ assuming an isothermal, optically-thin dust emission model
are the Boltzmann and Planck constants, respectively. The function \( f \) is defined as:

\[
  f = B(T_{\text{CMB}}) + A^{-1} \tilde{F}_{\nu},
\]

where \( B(T_{\text{CMB}}) = \exp[T_0/T_{\text{CMB}}] - 1 \), and \( T_{\text{CMB}} \) is the CMB temperature at a given redshift. The non-dimensional continuum flux \( \tilde{F}_{\nu} \) and the constant \( A \) correspond to:

\[
  \tilde{F}_{\nu} = 0.98 \times 10^{-16} \left( \frac{F_{\nu}}{\text{mJy}} \right),
\]

\[
  A = 4.33 \times 10^{-24} \left( \frac{g(z)}{g(6)} \right) \left( \frac{L_{\text{CII}}}{L_{\odot}} \right) \left( \frac{\alpha_{\text{CII}}}{M_{\odot}/L_{\odot}} \right) D,
\]

where \( g(z) = (1+z)/d_L^2 \), and \( d_L \) is the luminosity distance to redshift \( z \). Eq. 4 can be used to compute \( T_d \) using a single 1900 GHz observation (which provides both \( L_{\text{CII}} \) and \( F_{\nu} \)) modulo an estimate for \( D \) (eq. 3) and \( \alpha_{\text{CII}} \) (eq. 2).

Writing explicitly the expressions for \( D \) and \( \alpha_{\text{CII}} \) turns out that \( T_d \) is ultimately a function of \((\kappa, z, F_{\nu}, \Sigma_{\text{SFR}}, L_{\text{CII}}, y)\). All these parameters are generally constrained by observations out to very high redshift with two exceptions. Both \( \kappa \) and the metallicity \( Z \) are largely unknown. Hence in the derivation we consider a broad random uniform distribution for both parameters (see Section 3). To optimally constrain \( T_d \), we add the following broad physical conditions:

\[ \Box \] \( M_d \) cannot exceed the maximal dust production per supernova (SN), \( M_{d\text{max}} = 0.04 M_{\odot} \). This expression for \( M_{d\text{max}} \) is obtained by assuming a standard Salpeter 1 – 100 M_{\odot} Initial Mass Function (IMF, Ferrara & Tolstoy 2000), and that all the SN metal yield (\( \approx 2 M_{\odot} \) per SN) gets locked into dust grains (see eq. 15 in Sommovigo et al. 2021).

\[ \Box \] The IR-deduced star formation, \( SFR_{\text{IR}} = 10^{-10} L_{\text{IR}} \) (Kennicutt 1998), cannot largely exceed the SFR deduced from [C II] using the DL relation for starbursts\(^9\). We recall that the relation

\[
  L_{\text{IR}} = \left( \frac{M_d}{M_{\odot}} \right) \left( \frac{T_d}{8.5 \text{ K}} \right)^{4.5 \nu + 1} L_{\odot},
\]

holds for the dust model adopted here (Ferrara & al., in prep. 2022). Solutions not satisfying both these two bounds are discarded. These conditions result in a lower (upper) cut on \( T_d \) corresponding to unphysically large dust masses (FIR luminosities/SFR). This allows us to effectively constrain \( T_d \) at high-\( z \) despite of the lack of information on \((\kappa, Z)\).

This method for the dust temperature derivation has been tested on a sample of 19 local galaxies and four galaxies at \( z \approx 4 \) (Sommovigo et al. 2021; Baxx et al. 2021). For all these galaxies multiple data points in the FIR SED are available, allowing us to compare our inferred dust temperatures with robust \( T_d \) estimates obtained with traditional SED fitting. For the 19 local galaxies all the parameters \((\kappa, z, F_{\nu}, \Sigma_{\text{SFR}}, L_{\text{CII}}, y)\) are constrained by observations. This is also true for 3 out of the 4 high-\( z \) galaxies considered in the test\(^\Box\). For the remaining high-\( z \) galaxy, A1689-z1D at \( z = 7.13 \), both \((\kappa, Z)\) are unknown hence we assume the same broad distributions of values considered here for REBELS galaxies (\( Z = 0.3 - 1 Z_\odot \) and \( \kappa = 1 - 50 \)). These assumptions are motivated in detail in

\[ \Box \] Note that the value given here is 20\% lower w.r.t. the one given in Sommovigo et al. (2021) due to the slightly different dust model

\[ \Box \] Note that to set this broad upper limit we do not account for dust destruction.

\[ \Box \] Precisely, we allow SFR_{\text{IR}} to deviate at most by \( \pm 1 \) dex from SFR/(M_d \text{yr}^{-1}) = 10^{-7.06} \text{ L_{CII}}/L_\odot \) (De Looze et al. 2014), which is comparable to the dispersion around this relation observed at high-\( z \) (see e.g. Carniani et al. 2020)

\[ \Box \] For reference, in the considered local galaxies: 0.05 \( \leq Z/Z_\odot \leq 2.75 \), 0.1 \( \leq \kappa \leq 5.9 \) and [C II] emitting regions are as extended as stars (\( y \approx 1 \)). For the 3 high-\( z \) galaxies SP'TO418-47, MACS416-Y1 and B14-65666: 0.2 \( \leq Z/Z_\odot \leq 0.4 \), 9 \( \leq \kappa \leq 45 \) and 1.2 \( \leq y \leq 1.5 \).

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In Sec. 3, we recovered consistent dust temperatures with traditional SED fitting within 1σ spanning the very large redshift range $z = 0-8.31$ (as well as the large temperature range $20 \text{ K} \lesssim T_d \lesssim 100 \text{ K}$).

We also tested our method on simulations, applying it to the $z \sim 6.7$ galaxy Zinnia (a.k.a. serra05:s646:10643) from the SERRA simulation suite (Pallottini et al. 2019, see also Pallottini et al. 2022). Also in this case, we recover $T_d$ in agreement with single-temperature grey body SED fitting performed at the frequencies corresponding to ALMA bands 6, 7, and 8.

In (Sommovigo et al. 2021) we have shown that the dust temperature derived from our method matches that obtained from FIR SED fitting. We underline that this single value, which we refer to as $T_d$ throughout the paper, does not necessarily correspond to the dust physical temperature, which is instead characterised by a Probability Distribution Function (PDF, see e.g. Behrens et al. 2018; Sommovigo et al. 2020).

In general, $T_d$ does not necessarily provide a statistically sound representation of the PDF. We investigated the relation between the two in Appendix A of Sommovigo et al. (2021) for the simulated galaxy Zinnia. At $z \gtrsim 5$ such detailed comparison with observations is not currently possible due to the lack of information on the galaxies dust temperature PDF. In the case of Zinnia, we found that $T_d$ corresponds to the galaxy mass-weighted dust temperature ($\sim 60$ K), which is a factor $\sim 2$ colder than the luminosity-weighted dust temperature. In fact, Zinnia shows an excess of emission at MIR wavelengths (which is not traced by $T_d$) due to the presence of a scarce, but very luminous, hot dust component. The generalization of such PDF-$T_d$ relation is pending on a thorough comparison with a larger number of simulated galaxies, and future instruments observations possibly providing MIR-to-IR coverage for $z \gtrsim 5$ galaxies (such as the yet to be launched space telescopes Millimetron, see e.g. Novikov et al. 2021, and Origins, see e.g. Wiedner et al. 2021).

Finally, we remark that the increase by a factor $\sim 3$ in the stellar masses, which might result from using a non parametric SFH (Topping & al., in prep. 2022), barely affects our results. In fact, the derived values of $\alpha_{\text{CII}}$, $M_d$, and $T_{d,\text{IR}}$ of individual galaxies remain nearly unaffected, given their uncertainties, even by varying the stellar-mass based upper limits on $M_d$ by that amount.

### 3 APPLICATION TO REBELS GALAXIES

We use our method to derive $T_d$ for the 13 REBELS galaxies that are detected both in [C II] and in the continuum at 1900 GHz. In this work we always refer to this sub-sample of the survey targets. The observed properties of these sources are summarized in Table 1. For the data analysis we refer to the dedicated papers by Inami & al., in prep. 2022 (FIR continuum fluxes), Schouws & al., in prep. 2022 ([C II] luminosities), and Stefanon & al., in prep. 2022 (stellar masses).

For some of the REBELS sources the UV half-light radius has been measured by Bowler et al. (2017), who find on average $r_s = 1.3 \pm 0.8$ kpc. Bowler et al. (2017) caution that the large sizes $r_s > 1$ kpc are due to multiple components, with individual clumps being less extended $r_s \sim 0.2-1.0$ kpc. This is in agreement with the theoretical findings by Ferrara & al., in prep. (2022), who predicts sub-kpc sizes for most REBELS galaxies. In light of this, we consider star-forming/dust emitting regions of REBELS galaxies to be uniformly randomly distributed in the range $0.2 \text{ kpc} \lesssim r_s \lesssim 1.1 \text{ kpc}$.

We then consider the [C II] emitting regions to be $1.5 \lesssim \gamma \lesssim 3$ times more extended than the UV emitting ones, as indicated by previous observations of galaxies at $z \gtrsim 4$ (see e.g. Carniani et al. 2017, 2018; Matthee et al. 2017, 2019; Fujimoto et al. 2019, 2020; Ginolfi et al. 2020; Carniani et al. 2020).

The burstiness parameter of REBELS galaxies is unknown. High-$\gamma$ UV selected sources are expected to be highly star forming and UV emitting by construction (Dayal et al. 2013). Both locally and at intermediate redshift, values up to $\kappa_s \approx 100$ have been observed in such galaxies (see e.g. Daddi et al. 2010). Recently, Vallini et al. (2020) found for the mildly star-bursting COS-3018 at $z = 6.854$ a value of $\kappa_s \sim 3$, applying the [C II]-emission model given in Ferrara et al. (2019). Applying the same method on 11 bright UV selected galaxies\textsuperscript{11} at $z = 6-9$, Vallini et al. (2021) find values as large as $10 \lesssim \kappa_s \lesssim 80$. For the REBELS galaxies we choose a random uniform distribution in the range $1 \lesssim \kappa_s \lesssim 50$ inline with the independent derivation in Ferrara & al., in prep. (2022) which suggests that these sources are not extreme starbursts.

For the metallicity, which is also unconstrained by current data, we assume a broad uniform random distribution in the range $0.3-1$ Z$_\odot$. This general assumption is validated by numerical simulations of galaxies at $z \approx 6$ with similar stellar masses\textsuperscript{10} $< M_* < 10^{11}$ (Ma et al. 2016; Torrey et al. 2019), and several observational studies which analyse FIR lines (such as [N II], [N II], [C II], and [O III]) at $z \gtrsim 6-8$ to derive Z (see e.g. Pereira-Santaella et al. 2017; Hashimoto et al. 2019; De Breuck et al. 2019; Tamura et al. 2019; Vallini et al. 2020; Bakx et al. 2020; Jones et al. 2020, and references therein). Current estimates of Z at high redshift will be significantly ameliorated thanks to forthcoming ALMA observations and to the James Web Space Telescope (JWST) spectroscopy\textsuperscript{12}.

We pause for a quick summary: the physical properties needed for the application of our method are ($\kappa_s, z, P_{\text{m}}, Z, \Sigma_{\text{SFR}}, L_{\text{CII}}, y$). For REBELS galaxies all these properties are constrained by observations, with the exception of ($y, \kappa_s, Z$). For these three parameters, for each individual REBELS galaxy, we assume uniform random distributions in the broad ranges $1.5 \lesssim y \lesssim 3$, $1 \lesssim \kappa_s \lesssim 50$, and $0.3 < Z/Z_\odot < 1$, as motivated in the previous paragraphs.

Using eq. 2 we compute the [C II]-to-total gas conversion coefficient $\alpha_{\text{CII}}$ for all REBELS galaxies. Individual values are given in Table 1, the average is $\langle \alpha_{\text{CII}} \rangle = 5.3_{-1.8}^{+3.8}$. We note that this value is lower than the values we find in local galaxies (Sommovigo et al. 2021, $10 \lesssim \alpha_{\text{CII}} \lesssim 10^{3}$ at $z \approx 0$). Interestingly, we find low $\alpha_{\text{CII}} < 7$ also in the other $z > 4$ sources analysed in Sommovigo et al. 2021 (SPT0418-47 at $z = 4.225$, B14-65666 at $z = 7.15$, and MACS0416-Y1 at $z = 8.31$). This

\textsuperscript{11} In particular: MACS1149-JD1 (Hashimoto et al. 2018), A2744-YD4 (Laporte et al. 2017), MACS416- Y1 (Tamura et al. 2019; Bakla et al. 2020), SXDF-NB1006-2 (Inoue et al. 2016), B14-65666 (Hashimoto et al. 2019), BDF3299 (Carniani et al. 2017), J0121 J0253, J1211 (Harikane et al. 2020). See also Table B1, Fig. 3.

\textsuperscript{12} JWST will provide observations of optical nebular lines (such as H β, H α, [N II], [O III] and [O II]), which are reliable metallicity tracers, out to $z \approx 10$ (see e.g. Wright et al. 2010; Maiolino & Mannucci 2019; Chevallard et al. 2019)
might indicate that at fixed [C ii] luminosity high-z galaxies have a lower gas content (for a detailed discussion see e.g. Ferrara et al. 2019).

We conclude with a remark. The total gas-to-[C ii] conversion factor computed here $\alpha_{\text{CII}} = \Sigma_{\text{gas}}/\Sigma_{\text{CII}}$ is different by construction from the empirical molecular-to-[C ii] conversion factor $\alpha_{\text{CII, mol}} = \Sigma_{\text{H}_2}/\Sigma_{\text{CII}} = 31^{+16}_{-11}$ derived in Zanella et al. (2018). The ratio of the two conversion factors corresponds to $\alpha_{\text{CII, mol}}/\alpha_{\text{CII}} = f_{\text{H}_2} (r_{\text{gas}}/r_{\text{H}_2})^2 \approx f_{\text{H}_2} y^2$ (under the reasonable assumption that $r_{\text{CII}} \sim r_{\text{gas}}$ and $r_{\text{H}_2} \sim r_*$), where $f_{\text{H}_2} = M_{\text{H}_2}/M_{\text{gas}}$ is the molecular gas fraction. Since $f_{\text{H}_2} < 1$ by definition, the empirical result $\alpha_{\text{CII, mol}} \approx 31$ can be reconciled with the REBELS galaxies average value of $\langle \alpha_{\text{CII}} \rangle = 5.3$ for $y > 2.4$.

We proceed to compute $M_\odot$ and $T_d$ for all our targets. The results are summarised in Table 1, and discussed in detailed in the following dedicated Sections.

### 3.1 Dust temperatures

We find the median $T_d$ of REBELS galaxies to vary in the range between 39 K and 58 K with REBELS-39 (REBELS-18,32) being the galaxy with the warmest (coldest) dust. The median values are associated with large uncertainties (up to $\sim 35\%$) due to the lack of knowledge on the metallicity and burstiness of REBELS galaxies. Considering such uncertainties, the range widens to 30 K–73 K (where the lower limit is set by the condition on the dust masses and the upper one by that on the obscured SFRs, see Section 2). Note that these large uncertainties on $T_d$ are often comparable with the ones derived in the literature from traditional SED fitting using 2–3 data points at FIR wavelengths (see Table B1 where we compare $T_d$ values obtained from our method vs. SED fitting for several $z > 5$ galaxies). Further ALMA observations in bands 8–9 will help us significantly in constraining $T_d$ at high-z by sampling the SEDs closer to the FIR emission peak (see central panel of Fig. 1).

Overall we find an average dust temperature of $\langle T_d \rangle = 47 \pm 6$ K. This value is in agreement with the independent derivation in Ferrara & al., in prep. (2022), where they find $\langle T_d \rangle = (52 \pm 12)$ K. The physical model presented in Ferrara & al., in prep. (2022) relies on 1900 GHz continuum, and UV data (instead of [C ii] used here) to derive the physical properties of individual REBELS galaxies, such as...
their \((T_d, M_d, L_\text{IR}, \text{SFR})\), assuming assuming a given attenuation curve.

The average value \((T_d) \approx 47 \text{ K}\) is close to the dust temperature derived by Bethermin et al. (2020) from the mean stacked SED of ALPINE sources (and analogs in the COS-MOS field) in the redshift interval \(z = 5 - 6\) \((T_d = 43 \pm 5 \text{ K})\). For the lower redshift bin at \(z = 4 - 5\), where they have 3x more sources (analogats included), Bethermin et al. (2020) produce two different mean stacked SEDs dividing the considered sources depending on their SFR. For the galaxies with lower SFR (SFR > 100 \(M_\odot/\text{yr}\)) they find \(T_d = 41 \pm 1 \text{ K}\), while for the higher star forming galaxies (SFR > 100 \(M_\odot/\text{yr}\)) they find \(T_d = 47 \pm 2 \text{ K}\).

Since most REBELS galaxies have SFR \(\lesssim 100 \mathcal{M}_\odot/\text{yr}\) (both from the DL relation, and from the independent analysis of rest-frame UV and IR data performed in Ferrara \& al., in prep. 2022), they overall appear to have similar dust temperatures to the lower redshift ALPINE galaxies. We discuss this point in detail in the dedicated Section 4.

### 3.2 Dust masses

We find the median dust masses of REBELS sources to vary in the range 6.95 \(\lesssim \log(M_d/M_\odot) \lesssim 7.55\), REBELS-25 (REBELS-14,39) being the galaxy with the largest (lowest) dust mass. This is not surprising as REBELS-25 also has the largest stellar mass \(M_\star \sim 10^{10} \mathcal{M}_\odot\) and continuum emission \(F_{1000} = 260 \pm 22 \mu\text{Jy}\) in the sample. Overall we find an average dust mass around \((M_d) = (1.63 \pm 0.73) \times 10^7 \mathcal{M}_\odot\). This value is consistent with an independent derivation in Ferrara \& al., in prep. (2022) \((M_d) = (1.3 \pm 1.1) \times 10^7 \mathcal{M}_\odot\), although they find a larger scatter in the dust masses values.

We also compute the dust yield per SN, \(\gamma_d\), required to produce REBELS galaxies dust masses. We use that: \(\gamma_d = M_d/M_\star v_{\text{SN}}\), where \(v_{\text{SN}} = (53 \mathcal{M}_\odot)^{-1}\) is the number of SNe per solar mass of stars formed (Ferrara \& Tolstoy 2000) assuming a Salpeter 1 - 100 \(\mathcal{M}_\odot\) IMF. For all (but two of) the REBELS sources we find \(\gamma_d \lesssim 1 \mathcal{M}_\odot\) (see Table 1 for the value of \(\gamma_d\) in each galaxy), which is quite consistent with the latest constraints on SNe dust production by Lesniowska \& Michałowski (2019). They find that up to \(\gamma_d = 1.1 \mathcal{M}_\odot\) of dust per SN can be produced, whereas the exact value depends on the amount of dust which is destroyed/ejected during the SN explosion (1.1 \(\mathcal{M}_\odot\) corresponds to the case of no dust destruction/ejection).

We note that SN yield is still highly debated, with some works suggesting that dust destruction processes might only spare \(\lesssim 0.1 \mathcal{M}_\odot\) per SN (e.g., Bocchio et al. 2016; Matsuiura et al. 2019; Slavin et al. 2020). In this extreme case, interstellar medium grain growth or more exotic dust production mechanisms might be required (Mancini et al. 2015; Michałowski 2015; Graziani et al. 2020).

It is worth mentioning that the increase by a factor \(\sim 3\) of the stellar masses possibly resulting from a non parametric SFH prescription would have a relevant impact in this context. In fact, as explained in Sec. 2, the dust masses remain mostly unaffected. This implies, by definition, a reduction by the same factor \(\sim 3\) in the SN dust yield. As a result, we would find an average value of \((\gamma_d) \sim 0.23 \mathcal{M}_\odot\) for the REBELS sources, consistent even with stringent SN dust production constraints.

For a detailed discussion we refer to Dayal \& al., in prep. (2022) where they use the semi-analytical galaxy formation model DELPHI to investigate the dust build up and content of average \(z \sim 7\) galaxies. The processes accounted for include SNe dust production, grain growth, astraction, shock destruction, and ejection in outflows. The DELPHI model general predictions are then compared to the dust masses derived here for individual REBELS galaxies, finding a good agreement in most (77\%) cases.

### 3.3 IR emission

In Fig. 1 we show the FIR SEDs obtained for the REBELS galaxies considering a single-temperature grey body approximation and the dust masses and temperatures in Table 1. In the left panel we show the FIR SEDs obtained using the median \((T_d, M_d)\) values of each source. We can see that the flux at the peak of emission changes within a factor \(\times 10\), with a variation in the peak wavelength of emission 50 \(\mu\text{m} \lesssim \lambda_{\text{peak}} \lesssim 73 \mu\text{m}\).

Interestingly, almost all the \(\lambda_{\text{peak}}\) values that we predict lay within the range observable with ALMA, albeit in band-10 for the hottest sources.

For the FIR SED of each source we refer to Appendix A. Here we focus on REBELS-12 (central panel of Fig. 1), whose \(M_d\) and \(T_d\) are close to the average values in the sample. For REBELS-12 cold dust temperatures \(T_d < 30 \text{ K}\) are disfavoured as they would result in very large dust masses \(M_d > 10^5 \mathcal{M}_\odot\), and very low \(\kappa_\nu\) values that are unexpected for a strongly UV-emitting high-z source. Very hot dust temper-atures \((T_d > 80 \text{ K})\) are similarly unlikely as they would result in very large IR luminosities, and consequently unreasonably large obscured SFRs (for REBELS-12 \(SFR_\text{IR} > 500 \mathcal{M}_\odot/\text{yr}\)). These are hard to reconcile with the generally blue UV slopes \(\beta\) observed in most REBELS galaxies (e.g. \(\beta = -1.99\) for REBELS-12). We briefly discuss a possible cave in Section 5.

3.3.1 \(F_{1000}\) to total IR luminosity conversion at \(z \sim 7\)

Using eq. 7 we can compute the IR luminosities of all the REBELS galaxies. We find that their IR luminosities vary in the range \(1.7 \times 10^{11} \mathcal{L}_\odot \lesssim L_\text{IR} \lesssim 2.8 \times 10^{12} \mathcal{L}_\odot\), which corresponds to obscured SFRs in the range \(17 - 285 \mathcal{M}_\odot/\text{yr}\) (assuming the conversion factor given in Table 1). Interestingly, the most IR luminous REBELS (REBELS-25) classifies as an Ultra-Luminous InfraRed Galaxy (ULIRG, with \(L_\text{IR} > 10^{12} \mathcal{L}_\odot\), see e.g. Lonsdale et al. 2006) despite being UV selected. This peculiar source will be discussed in detail in Hygate \& al., in prep. (2022).

Combining these values with the total star formation rates derived using the DL relation for starbursts we find that on average \((SFR_\text{IR}/SFR) = 53\%\) (in agreement with the independent derivation in Schouws et al. 2021; Ferrara \& al., in prep. 2022; Dayal \& al., in prep. 2022), implying that REBELS galaxies are on average relatively UV-transparent sources. The only exceptions are represented by REBELS-8, REBELS-14, REBELS-25 and REBELS-39 for which we find \(SFR_\text{IR} \lesssim SFR\), i.e. the SFR deduced from \([\text{C}\text{II}]\) is exceeded (barely, only by 6\%), in the case of REBELS-8. Interestingly these sources include 2 (REBELS-8, and the peculiar REBELS-25) of the 4 galaxies for which the UV-to-IR emission model in Ferrara \& al., in prep. (2022) fails to find a solution for \((T_d, M_d)\) (using the same dust model adopted here,
The monochromatic luminosity at ν as:

\[ L_{\nu} \]

indicate that their UV and IR emitting regions are spatially to-UV flux ratios compared to their UV slopes. This might i.e. MW-like dust). In fact these galaxies have very large IR-to-UV flux ratios compared to their UV slopes. This might be impossible to constrain their \( T_d \) with ordinary SED-fitting.

More recently, \( T_d \) obtained from stacked SEDs at \( z < 6 \) has been combined with the \( T_d \) derived for individual galaxies detected in multiple bands at \( z \lesssim 5 \). Two studies reached somewhat discordant conclusions. While Liang et al. (2019); Pässler et al. (2020) deduce a flattening of \( T_d - z \) relation at \( z > 4 \), Bouwens et al. (2020) confirm the trend with little modification in the slope.

Either trends are yet to be explained by a physical model. Previous theoretical works suggested that warmer dust temperatures at high-\( z \) might depend on the dust geometry, with most of the dust in high-\( z \) galaxies being located in compact, young star-forming regions, where the dust heating is particularly efficient (Behrens et al. 2018; Liang et al. 2019; Sommovigo et al. 2020). Other works (Magnelli et al. 2014; Ma et al. 2016, 2019; Shen et al. 2021) have suggested that the increasing \( T_d(z) \) trend together with the flattening at \( z \sim 4 \) are correlated to the specific star formation rate increase with redshift and its subsequent plateau at \( z \sim 4 \) (see e.g. Tomczak et al. 2016; Santini et al. 2017). Finally the lower dust-to-metal ratio predicted for early galaxies from simulations at \( z \gtrsim 5 \) (see e.g. Aoyama et al. 2017; Behrens et al. 2018; Shen et al. 2021; Pallottini et al. 2022) could also motivate the presence of hot dust at high-\( z \). In fact, a lower dust content results in warmer dust temperatures for a given \( L_{\text{FIR}} \).

Our analysis of the REBELS sample can clarify the issue of the \( T_d(z) \) evolution in a unique way for two reasons: (a) REBELS continuum detections double the number of previously known sources at \( z \lesssim 5 \); (b) possible deviations from the linear trend are more robustly identified at the higher (\( z \approx 7 \)) redshift of the REBELS sample.

Extrapolating the Schreiber et al. (2018) relation to \( z = 7 \) gives \( T_d = (56 \pm 4) \) K, i.e. a slightly warmer temperature than found here, \( T_d \sim 47 \) K. This does not come as a surprise, as there is no special physical motivation to expect a linearly increasing \( T_d \) trend. It is therefore crucial to develop a simple but physical theoretical framework against which observations of individual UV-to-FIR detected galaxies in the range \( 0 \lesssim z \lesssim 8.5 \) can be compared and interpreted. This is our goal in the next Section.

### 4 DUST TEMPERATURE EVOLUTION

In the last few years, several works have suggested the presence of a trend of increasing dust temperature with redshift in star forming galaxies detected in the UV-to-IR rest-frame wavelength range (Magdis et al. 2012b; Magnelli et al. 2013, 2014; Béthermin et al. 2015; Schreiber et al. 2018).

In particular, Schreiber et al. (2018) fitted stacked SEDs to a complete sample of main sequence galaxies in the CANDELS fields, and derived the linear relation \( T_d = 32.9 + 4.6(z-2) \), shown in the inset of Fig. 3. Compared to those for individual sources, stacked SEDs (a) reduce the scatter introduced in the relation by different intrinsic galaxy properties, and (b) allow to extend the relation up to \( z \sim 4 \). The above trend is also consistent with the stacked SED fitting results produced for ALPINE galaxies by Béthermin et al. (2020) at \( 4 < z < 6 \) (also shown in the inset). In this case, a stacking procedure was necessary since ALPINE galaxies are individually detected only in a single band at restframe 1900 GHz. Hence it would be impossible to constrain their \( T_d \) with ordinary SED-fitting.

4.1 Physical origin of the \( T_d - z \) relation

Under the assumption that FIR and UV emitting regions are co-spatial, \( L_{\text{IR}} \) is equal to the absorbed fraction of the intrinsic UV luminosity, \( L_{1500} \):

\[ L_{\text{IR}} = (1 - e^{-\tau_{d}})L_{1500} = (1 - e^{-\tau_{d}})K_{1500}SFR. \] (9)

\[ L_{\nu} = g(\nu) \frac{F_{\nu}}{m_{\text{Jy}}} L_\odot \] (8)

where \( g(z) \) is defined in eq. 6. The results for the conversion factor \( L_{\text{IR}}/v_0L_{\nu_0} \) are shown in Fig. 2.

We find a median value of \( L_{\text{IR}}/v_0L_{\nu_0} = 14^{+8}_{-5} \) (where both \( L_{\text{IR}}, v_0L_{\nu_0} \) are in solar luminosity units). This range is consistent with the one obtained in Bouwens et al. (2020) at \( z \sim 7 \) using a simple modified black body with \( \beta_\nu = 1.6 \), and a dust temperature derived from a linear fit of currently available dust temperature data vs. redshift. It is worth noting that they consider \( T_d = 54 \) K at \( z = 7 \), which is warmer than the temperatures derived here for the REBELS galaxies. We will discuss this point in detail in the following Section 4.
In the previous equation, \( \tau_{\text{eff}} \) is the galaxy effective dust attenuation optical depth at 1500 \( \AA \), which directly relates to the transmissivity, \( \tau_{\text{eff}} = -\ln T \). The transmissivity is defined as the ratio of the observed-to-intrinsic UV luminosity, i.e. \( T = 1 \) (\( T = 0 \)) for a fully transparent (obscured) galaxy. Note that, depending on the relative dust and star distributions, \( \tau_{\text{eff}} \) might significantly differ from the physical UV optical depth, \( \tau_{1500} \). Hence, a large transmissivity does not directly imply a low physical optical depth. The conversion factor for (a) continuous star formation at a fix age of 150 Myr, (b) Salpeter IMF in the range 1-100 \( M_\odot \), (c) metallicity \( Z = 1/3 \) Z\( _\odot \) is \( k_{1500} \equiv L_{1500}/SFR = 1.174 \times 10^{10} L_\odot/(M_\odot \text{yr}^{-1}) \) (Ferrara & al., in prep. 2022).

Combining eq. 7 and 9, and recalling that \( M_d = DM_\odot \), it follows that

\[
T_d = 29.7 \left( \frac{1 - e^{-\tau_{\text{eff}}}}{Z} \right)^{1/(4+\beta_d)} \left( \frac{\text{Gyr}}{\frac{t_{\text{dep}}}{10^9}} \right) \text{K}
\tag{10}
\]

where \( \beta_d = 2.03 \) and \( Z \) is in solar units. We have introduced the total gas depletion time \( t_{\text{dep}} = M_J/SFR \), which we derive in the following from cosmological arguments. Eq. 10 shows that the dust temperature is larger in optically thick, low metallicity systems with a short depletion time.

Let us express \( M_b \) and SFR of a galaxy in terms of its total (dark + baryonic) halo mass, \( M \), and mean dark matter accretion rate, \( (dM/dt) \):

\[
M_b = f_b M - M_*
\tag{11}
\]

\[
\text{SFR} = \epsilon_* f_b \left( \frac{dM}{dt} \right)
\tag{12}
\]

where \( \epsilon_* \) is the star formation efficiency, and \( f_b = \Omega_b/\Omega_m \) is the (cosmological) baryon fraction in the halo. By integrating eq. 12 and substituting it into eq. 11 we obtain

\[
M_b = f_b M (1 - \epsilon_*)
\tag{13}
\]

Numerical simulations (Fakhouri et al. 2010; Dekel & Krumholz 2013; Correa et al. 2015) provide the following fit to the mean halo accretion rate, as a function of redshift and halo mass, \( M_{12} = M/10^{12} M_\odot \),

\[
\left( \frac{dM}{dt} \right) = 69.3 M_{12} f'(z) E(z) M_\odot \text{yr}^{-1}.
\tag{14}
\]
with
\[ f'(z) = -0.24 + 0.75(1 + z); \quad E(z) = \left( \Omega_m (1 + z)^3 + \Omega_{\Lambda} \right)^{1/2}. \]
By combining the previous equations \( t_{\text{dep}} \) takes the form
\[ t_{\text{dep}}(z) = t_{\text{dep}}(f'(z)E(z))^{-1}, \]
where the timescale \( t_{\text{dep}} = 14.4(1 - \epsilon_*)/\epsilon_* \) Gyr is fixed so that \( t_{\text{dep}}(z = 0) = 2 \) Gyr as approximately measured in local galaxies, including the MW (Bigiel et al. 2008; Leroy et al. 2008; Genzel et al. 2010).

According to eq. 16, the depletion time decreases with redshift as \((1 + z)^{-1/2}\) as a result of the higher cosmological accretion rate at early times. This point is crucial as, barring variations of the optical depth and metallicity (see below), the redshift evolution of \( T_d \) is governed by the gas depletion time in galaxies. From the result above, and using eq. 10, it follows that
\[ T_d \propto (1 + z)^{5(2+z)(\delta s)} \approx (1 + z)^{0.42}. \]

As we will see shortly, this trend matches perfectly the observational trend for the rest of the universe. At fixed redshift, scatter is introduced by variations in dust temperatures, at fixed redshift as \( \delta s = 0 \) Gyr as approximately measured in local galaxies, including the MW (Bigiel et al. 2008; Leroy et al. 2008; Genzel et al. 2010).

4.2 Comparison with observations at \( 0 < z < 8 \)

We now intend to compare our theoretical predictions with dust temperature estimates available in the literature for UV-detected sources at \( 0 < z < 8 \).

We recover the dust temperatures of individual UV-to-IR detected galaxies whose stacked SEDs in the redshift range \( 0 < z < 3 \) are used in the analysis by Schreiber et al. (2018). We then add all the UV-selected galaxies at \( z < 5 \) for which dust temperature estimates are available in the literature (see Table B1 for details of the sources). We apply the method used here to derive \( T_d \) for these galaxies, finding values consistent (within \( 1 - \sigma \) with SED fitting results (see Table B1 for the detailed comparison). Finally, we apply our method to individual ALPINE galaxies detected simultaneously in [C ii] and continuum. We find their median dust temperatures to vary in the range \( 35 \leq T_d \leq 60 \) K, which is consistent with the stacked SEDs fitting results in Bethermin et al. (2020) (40 K \( \leq T_d \leq 49 \) K). A detailed analysis of ALPINE galaxies will be presented in Sommovigo et al., in prep. (2022).

The complete collection of \( T_d \) values is shown in Fig. 3 as a function of redshift. We stress that we consistently compare dust temperatures obtained by fitting individual galaxy SEDs; moreover, the same method is applied to all high-z sources (see Appendix B). This avoids the confusion arising from comparing intrinsically different quantities such as dust temperatures obtained from stacked SEDs, and/or peak dust temperatures \( T_{\text{peak}} \sim 2.9 \times 10^{2}(\alpha_{\text{peak}}/\mu)^{-1} \).

The physical interpretation of \( T_{\text{peak}} \) might be unclear for \( z > 5 \) galaxies. Indeed currently available data at these redshifts hardly trace the peak of FIR emission. Moreover, when a different SED fitting function other than the optically thin grey-body is used, \( T_{\text{peak}} \) can significantly differ from \( T_d \). In fact the assumptions made for the MIR (rest-frame) portion of the spectra affect \( T_{\text{peak}} \) (Faisst et al. 2020), and the validity of such assumptions cannot be tested as no currently available instrument probes MIR wavelengths at \( z > 5 \).

We find that our predictions are in agreement with data. Fitting all the dust temperatures with a single power law: \( T_d(z) = ax^b + f \), we find \( a = (0.58 \pm 0.04) \), which is close to the value 0.42 given in eq. 17. The slight difference is due to the fact that \( T_d \) does not depend uniquely on redshift, as discussed in detail in the previous section (see eq. 18). Hence fitting all the data with a single power-law is misleading.

The additional dependence on the column density (for optically thin sources), and metallicity (for optically thick sources) is responsible for the scatter in the measured temperatures at a given redshift. At \( z \approx 0.3 \variations as large as \( \Delta T_d \approx 22 \) K are observed (Schreiber et al. 2018), which is perfectly consistent with our predictions (\( \Delta T_d \approx 25 \) K in the local Universe).

The amplification of the dust temperature scatter at high-z that we predict (if the \( N_H, Z \) range does not evolve) is also consistent with data. In the narrow redshift range \( 7.6 < z < 8.3 \)
variations as large as ΔT_d = 53 K are observed (we predicted ΔT_d ≈ 55 K at z = 8).

At one extreme there are galaxies hosting very hot dust such as MACS0416-Y1 and A2744-YD4 (T_d > 80 K Bakx et al. 2020, and T_d = 90 ± 20 K Laporte et al. 2017; Behrens et al. 2018, respectively). On the other, there are galaxies such as REBELS-18 and J1211-0118 with lower effective optical depths τ_eff < 1, i.e. larger UV transmissivity (ultimately resulting in less efficient dust heating). For these galaxies we predict mean gas column densities around 0.3 × 10^{20} cm^{-2} ≲ N_H ≲ 1.0 × 10^{20} cm^{-2}.

5 SUMMARY AND CONCLUSIONS

We have applied a novel method (Sommovigo et al. 2021) to derive the dust temperature to 13 z ≈ 7 galaxies part of the ALMA Large Program REBELS. Our method combines the continuum and super-imposed [C II] line emission measurements, thus breaking the SED fitting degeneracy between dust mass and temperature. This allows us to constrain T_d from a single-band restframe observation at 1900 GHz, and to derive dust masses, IR luminosities, and the obscured SFR. Moreover, since REBELS targets constitute the first significant sample of continuum detected sources at z ~ 7 (for which T_d estimates are available), we can extend the reported T_d-redshift relation (Magdis et al. 2012a; Magnelli et al. 2013; Béthermin et al. 2015; Schreiber et al. 2018; Faisst et al. 2020; Bouwens et al. 2020) into the Epoch of Reionization.

We summarize below our main findings:

- **Dust temperature and mass**: the median T_d values for REBELS galaxies vary in the range 39 – 58 K, with ~ 35% associated uncertainty. The median dust masses are in the narrow range (0.9 – 3.6) × 10^{9} M_{⊙}. Dust production from SNe alone in most cases (85%) can generate such dust masses assuming a dust yield $\lesssim 1$ M_{⊙} per SN;

- **IR luminosities and L_{1000}-to-L_{IR} conversion at z ~ 7**: REBELS galaxies IR luminosities vary in the range 1.7 × 10^{11} L_{⊙} ≲ L_{IR} ≲ 2.8 × 10^{12} L_{⊙}, which corresponds to obscured SFRs around $\sim 17 - 285$ M_{⊙}/yr. We also derive their average conversion factor $L_{IR} = 14^{+3}_{-5}$ L_{1000}, where $L_{1000} = \nu_0 L_0$, with $\nu_0 = 1900$ GHz. This value is consistent with an extrapolation of the empirical fitting formula of Bouwens et al. (2020) to z ≈ 7;

- **Dust temperature cosmic evolution**: we produce a new physical model (see eq. 10-18) that motivates the dust temperature increase with redshift. Such trend is an imprint of the decreasing gas depletion time towards high-z, $t_{dep} \propto (1 + z)^{1/2}$. We show that $T_d \propto t_{dep}^{-1/6}$, or $T_d \propto (1 + z)^{0.42};$

- **Dust temperature scatter at a given redshift**: on top of the T_d-z trend, we can also physically motivate the scatter in the measured T_d values at a given redshift. We find that in UV-transparent galaxies (UV transmissivity $\gtrsim 37\%$) the scatter in T_d depends solely on the column density N_H, with larger N_H corresponding to hotter dust. Instead, in UV-obscured galaxies the scatter in T_d depends only on the metallicity Z, with lower Z implying hotter dust.

A very hot dust component, implying a large obscured SFR, can coexist with a steep UV slope in the presence of **spatial segregation** of IR and UV emitting regions. This possibility has been suggested by theoretical studies and simulations in some z ~ 7 – 8 sources (Behrens et al. 2018; Liang et al. 2019; Sommovigo et al. 2020). Such scenario is also supported by some observations; for instance Hodge et al. (2012, 2016); Carniani et al. (2017); Laporte et al. (2017); Bowler et al. (2018, 2021) find significant spatial offset between their ALMA and HST data. High-resolution ALMA follow-up observations of REBELS galaxies are required in order to make a step forward (see also Inami et al., in prep. 2022; Ferrara & et al., in prep. 2022 for a discussion on this point).

JWST will also provide us with much more accurate metallicity measurements at z ~ 5, improving current estimates of the dust-to-gas ratios at high-z. Finally, further ALMA observation at shorter wavelengths in band 7 – 8 – 9, will allow us to reduce the uncertainties in current dust temperatures estimates at z ~ 5 by sampling galaxies SEDs closer to the FIR emission peak (see Fig. 1 and the discussion in Bakx et al. 2021).

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DATA AVAILABILITY

Data generated in this research will be shared on reasonable request to the corresponding author.

REFERENCES

Aoyama S., Hou K.-C., Shimizu I., Hirashita H., Todoroki K., Choi J.-H., Nagamine K., 2017, MNRAS, 466, 105
Atek H., et al., 2015, Astrophysical Journal, 814
Balzi T. J. L. C., et al., 2020, MNRAS, 493, 4294
Balzi T. J. L. C., et al., 2021, arXiv e-prints, p. arXiv:2108.13479
Barisic I., et al., 2017, The Astrophysical Journal, 845, 41
Behrens C., Pallottini A., Ferrara A., Gallerani S., Vallini L., 2018, MNRAS, 477, 552
Bethermin M., et al., 2015, A&A, 573, A113
Bethermin M., et al., 2020, arXiv e-prints, p. arXiv:2002.00962
Bigiel F., Leroy A., Walter F., Brinks E., de Blok W. J. G., Madore B., Thornley M. D., 2008, AJ, 136, 2846
Blain A. W., Barnard V. E., Chapman S. C., 2003, Monthly Notices of the Royal Astronomical Society, 338, 733
Bocchio M., Marassi S., Schneider R., Bianchi S., Limongi M., Chieffi A., 2016, A&A, 587, A157
Bouwens R. J., et al., 2010, ApJL, 709, L133
Bouwens R. J., et al., 2011, Nature, 469, 504
Bouwens R. J., et al., 2015, ApJ, 803, 34
Bouwens R. J., et al., 2016, ApJ, 833, 72
Bouwens R., et al., 2020, ApJ, 902, 112
Bouwens R. J., et al., 2021, arXiv e-prints, p. arXiv:2106.13719
Bowler R. A. A., et al., 2015, MNRAS, 452, 1817
Bowler R. A. A., Dunlop J. S., McLure R. J., McLeod D. J., 2017, MNRAS, 466, 3612
Bowler R. A. A., Bourne N., Dunlop J. S., McLure R. J., McLeod D. J., 2018, MNRAS, 481, 1631
Bowler R. A. A., Cullen F., McLure R. J., Dunlop J. S., Avison A., 2021, MNRAS, 41, 1201
Bradley L. D., et al., 2012, ApJ, 760, 11
Bradley L. D., et al., 2014, ApJ, 792, 76
Calzetti D., Armus L., Bohlin R. C., Kinney A. L., Koornneef J., Storchi-Bergmann T., 2000, The Astrophysical Journal, 533, 682
Capak P., et al., 2015, Nature, 522, 455
Carilli C. L., et al., 2013, ARA&A, 51, 105
Carilli C. L., et al., 2016, The Astrophysical Journal, 833, 73
Carniani S., et al., 2017, A&A, 605, A42
Carniani S., et al., 2018, MNRAS, 478, 1170
Carniani S., et al., 2020, arXiv e-prints, p. arXiv:2006.09402
Casey C. M., 2012, MNRAS, 425, 3094
Casey C. M., Hodges J., Zavala J. A., Spilker J., da Cunha E., Staguhn J., Finkelstein S. L., Drew P., 2018, The Astrophysical Journal, 862, 78
Chabrier G., 2003, ApJL, 586, L133
Chevallard J., Charlot S., 2016, Monthly Notices of the Royal Astronomical Society, 462, 1415
Chevallard J., et al., 2019, MNRAS, 483, 2621
Correa C. A., Wyithe J. S. B., Schaye J., Duffy A. R., 2015, MNRAS, 450, 1521
Daddi E., et al., 2010, The Astrophysical Journal, 714, L118
Dale D. A., et al., 2012, The Astrophysical Journal, 745, 95
Dayal P., Ferrara A., 2018, Phys. Rep., 780, 1
Dayal P., et al., in prep. 2022, 0, 0
Dayal P., Dunlop J. S., Maio U., Ciardi B., 2013, MNRAS, 434, 1486
De Breuck C., et al., 2019, A&A, 631, A167
De Looze L., et al., 2014, A&A, 568, A62
Dekel A., Kruijssen J. M. D., 2013, Monthly Notices of the Royal Astronomical Society, 432, 455
Draine B., 1989, Infrared spectroscopy in astronomy.
Draine B., 2003, Annual Review of Astronomy and Astrophysics, 41, 241
Draine B. T., Li A., 2007, ApJ, 657, 810
Dunlop J. S., 2016, The Messenger, 166, 48
Dunlop J. S., et al., 2013, Monthly Notices of the Royal Astronomical Society, 432, 3520
Dunne L., Eales S. A., 2001, Monthly Notices of the Royal Astronomical Society, 327, 697
Ellis R. S., et al., 2012, The Astrophysical Journal, 763, L7
Faiss A. L., Fudamoto Y., Oesch P. A., Scoville N., Riechers D. A., Pavesi R., Capak P., 2020, arXiv e-prints, p. arXiv:2005.07716
Fakhouri O., Ma C.-P., Boylan-Kolchin M., 2010, Monthly Notices of the Royal Astronomical Society, 406, 2267
Ferrara A., Tolstoy E., 2000, Monthly Notices of the Royal Astronomical Society, 313, 291
Ferrara A., et al., in prep. 2022, 0, 0
Ferrara A., Vallini L., Pallottini A., Gallerani S., Carniani S., Kobulnicky M., Decataldo D., Behrens C., 2019, Monthly Notices of the Royal Astronomical Society, 489, 1
Fudamoto Y., et al., 2020, MNRAS, 491, 4724
Fudamoto Y., et al., 2021, Nature, 597, 489–492
Fujimoto S., et al., 2019, ApJ, 887, 107
Fujimoto S., et al., 2020, arXiv e-prints, p. arXiv:2003.00013
Galametz M., et al., 2012, Monthly Notices of the Royal Astronomical Society, 425, 763
Galliano F., Dwek E., Chiang P., 2008, ApJ, 672, 214
Genzel R., et al., 2010, MNRAS, 407, 2001
Ginolfi M., et al., 2020, A&A, 633, A90
Graziani L., Schneider R., Ginolfi M., Hunt L. K., Maio U., Glatzle M., Ciardi B., 2020, MNRAS, 494, 1071
Harikane Y., et al., 2020, ApJ, 896, 93
Hashimoto T., et al., 2018, Nature, 557, 392
Hashimoto T., et al., 2019, Pub. Astron. Soc. Japan, 71, 71
Heiderman A., Evans N. J., Allen L. E., Huard T., Heyer M., 2010, The Astrophysical Journal, 723, 1019
Hildebrand R. H., 1983, QJRAS, 24, 267
Hodge J. A., da Cunha E., 2020, Royal Society Open Science, 7, 200556
Hodge J. A., Carilli C. L., Walter F., de Blok W. J. G., Riechers D., Daddi E., Lentati L., 2012, ApJ, 760, 11
Hodge J. A., et al., 2016, ApJ, 833, 103
Hollenbach D. J., Tielens A. G. G. M., 1999, Rev. Mod. Phys., 71, 173
Hygate A., et al., in prep. 2022, 0, 0
Inami H., al., in prep. 2022, 0, 0
Inoue A. K., et al., 2016, Science, 352, 1559
Inoue A. K., Hashimoto T., Chihara H., Koike C., 2020, MNRAS, 335, 753
Jones T., Sanders R., Roberts-Borsani G., Ellis R. S., Laporte N., Treu T., Harikane Y., 2020, arXiv e-prints, p. arXiv:2006.02447
APPENDIX A: INDIVIDUAL SEDS

In Fig. B1 we show the individual FIR SEDs obtained for all REBELS [C ii] and continuum detected galaxies (analogously to the central panel in Fig. 1 representing REBELS-12 only). For all the sources we show the SEDs obtained with their median values of \((T_d, M_d)\), and the variation due to the 1-\(\sigma\) uncertainties in these two quantities (on average \(\Delta T_d/T_d \sim 30\%\) and \(\Delta M_d/M_d \sim 70\%\)).

These large uncertainties result from the lack of information on the metallicity Z and burstiness parameter \(\kappa_s\) of REBELS galaxies (see the discussion in Section 3). Future observations will help us to constrain both these quantities, thus reducing the uncertainties in the predicted \(T_d, M_d\). In fact, with ALMA we can investigate the [O III]/[C ii] luminosities ratios of REBELS sources, which can be used to reliably constrain their \(\kappa_s\) using the model in Vallini et al. 2021 (these ALMA observations would also provide us with an additional data point in the FIR continuum underlying [O iii]). Moreover, future JWST optical nebular lines observations will allow us to improve metallicity estimates out to very high-z, possibly reaching a precision as low as \(\Delta Z/Z \sim 35\%\) at \(z \sim 7\) (Wright et al. 2010; Chevallard et al. 2019; Maiolino & Mannucci 2019).

APPENDIX B: APPLICATION TO OTHER SAMPLES

In Table B1 we show the comparison between measured dust temperatures for galaxies at \(z > 5\) available in the literature compared with the results from our method. Here we briefly discuss the assumptions used in our derivation.

For the galaxies MACS0416-Y1, B14-65666 and A1689-zD1 we refer to the dedicated works discussing the application of our method to multiple-band SED fitting (MACS0416-Y1, B14-65666: Sommovigo et al. 2021, A1689-zD1: Bakx et al. 2021). We highlight that A1689-zD1 is the only \(z > 5\) galaxy for which a band-9 continuum detection, short-wards of the peak of FIR emission, is available. It is very promising that also in this case -where traditional SED fitting is particularly precise thanks to the widespread continuum data available- our method gives a consistent \(T_d\) value (within \(\sim 0.5\) \(\sigma\)).

For the remaining sources detected in both [O iii]88\(\mu m\) and [C ii]158 \(\mu m\) (all but HZ4-10), Vallini et al. (2021) derived the value of \((\kappa_s, Z)\), albeit with large uncertainties (on average: \(10 \leq \kappa_s \leq 80\) and \(0.2 \leq Z/Z_\odot \leq 0.4\); see Table 1 in the paper). For each individual source we assume a random distribution for \((\kappa_s, Z)\) around the mean value, with the dispersion corresponding to the uncertainty. For all of these sources the ratio \(y = r_{\text{C I}}/r_{\text{sil}}\) is also measured: \(y = 1.8 \pm 0.4\) (J1211-0118 and J0217-0208, Harikane et al. 2020), \(y = 8 \pm 2\) (A2744-YD4, Laporte et al. 2019). Finally, for HZ4-10, due to the lack of observational constraints on \((Z, \kappa_s)\), we rely on the same assumptions used for both the ALPINE individual galaxies and the REBELS galaxies, which are extensively described and motivated in Section 3.

In all cases with our method we derive \(T_d\) values consistent with literature estimates well within \(\pm 30\%\). The more discrepant case is represented by J0217-0208 for which we predict a warmer temperature, albeit consistent within the large uncertainty given from SED fitting (\(\sim 52\%\)).

We note that for this galaxy, assuming the dust model adopted here, from traditional SED fitting (relying only on the two detections at 120\(\mu m\), 158\(\mu m\)) one would deduce a very large dust mass \(\log(M_d/M_\odot) = 9.28 \pm 0.17\). Given the stellar mass estimated for this galaxy \((M_* \sim 3 \times 10^9 M_\odot\) from Harikane et al. 2020), such dust mass would imply a very large dust yield of \(\gamma_d = 3 \times M_\odot\) per SN, which is not compatible with SNe dust production constraints (Bocchio et al. 2016; Matsuura et al. 2019; Le´ sniewska & Micha lowski 2019; Slavin et al. 2020). Further ALMA observations at shorter wavelengths will help us understand weather the dust temperature of this galaxy has been underestimated (and thus the dust mass overestimated), which would reduce the tension with the \(T_d\) value derived with our method. An other possibility is that the stellar mass of this source has been underestimated, this would relax the requirements set by the condition on \(M_d < M_{d,\text{max}} \propto M_{*\text{sil}}\) (see Section 2), allowing for larger dust masses and lower temperatures in our derivation. If neither of these possibilities is verified, alternative dust production scenarios might have to be invoked for this source (see the discussion in Section 3.2).

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\begin{table}[ht]
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\textbf{Derived} & \textbf{ID\#} & \textbf{Literature} \\
\hline
\multicolumn{3}{l}{Table B1. Measured dust temperatures for galaxies at \(z > 5\) available in the literature compared with the results from our method. We note that for the literature data, we always show the result derived from traditional SED fitting in the most recent reference. The only case in which we provide two estimates is that of A2744-YD4, as the latest one is obtained from ad hoc simulations rather than direct measurements (2, Behrens et al. 2018). These temperatures are the ones shown in Fig. 3 as grey triangles (literature data) and stars (our derivations). References: 1 (Laporte et al. 2019), 2 (Behrens et al. 2018), 3 (Bakx et al. 2020), 4 (Tamura et al. 2019), 5 (Hashimoto et al. 2019), 6 (Bowler et al. 2018), 7 (Bakx et al. 2021), 8 (Knudsen et al. 2016b), 9 (Watson et al. 2015), 10 (Harikane et al. 2020), 11 (Faist et al. 2020), 12 (Pavesi et al. 2016b; Pavesi et al. 2019), 13 (Capak et al. 2015).} \\
\hline
\end{tabular}
\end{table}

\textsuperscript{16} The minor differences in the quoted \(T_d\) values arise form the change in the adopted dust model, see also Section 2.
Figure B1. Variation in the derived SEDs of all REBELS [C II] and continuum detected galaxies due to the 1 − σ uncertainties in their individual ($T_d, M_d$). The SEDs are colour coded according to the corresponding dust temperatures (see colorbar). The dashed black curves show the SEDs obtained with the median ($T_d, M_d$) values for each galaxy. The black points represent the continuum observations at 1900 GHz. For further details on the sources see Table 1.)