Towards an understanding of the Of?p star HD 191612: optical spectroscopy

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ABSTRACT

We present extensive optical spectroscopy of the early-type magnetic star HD 191612 (O6.5f?pe–O8fp). The Balmer and He I lines show strongly variable emission which is highly reproducible on a well-determined 538-d period. He II absorptions and metal lines (including many selective emission lines but excluding He II \( \lambda 4686 \) Å emission) are essentially constant in line strength, but are variable in velocity, establishing a double-lined binary orbit with \( P_{\text{orb}} = 1542 \) d, \( e = 0.45 \). We conduct a model-atmosphere analysis of the spectrum, and find that the system is consistent with a \( \sim \)O8 giant with a \( \sim \)B1 main-sequence secondary. Since the periodic 538-d changes are unrelated to orbital motion, rotational modulation of a magnetically constrained plasma is strongly favoured as the most likely underlying ‘clock’. An upper limit on the equatorial rotation is consistent with this hypothesis, but is too weak to provide a strong constraint.

Key words: stars: binaries: spectroscopic – stars: early-type – stars: individual: HD 191612 – stars: magnetic fields – stars: rotation.

1 INTRODUCTION

Peculiarities in the spectrum of the early-type star HD 191612 were first noted by Walborn (1973); it remains one of only three known Galactic examples of the Of?p class,\(^1\) this designation indicating C III \( \lambda 4650 \) in emission with comparable strength to N III \( \lambda 4640 \). Renewed interest followed the discovery of recurrent spectral variabil-

\(^1\) The others are HD 108 and HD 148937; the Small Magellanic Cloud Of?p stars AzV 220 and 2dFS 936 were found subsequently (Walborn et al. 2000; Massey & Duffy 2001; Evans et al. 2004).
ity between spectral types O6–O7 and O8 (with correlated changes in the unusual emission-line features; Walborn et al. 2003), which Walborn et al. (2004) showed to be consistent with a \sim 540 d period identified in Hipparcos photometry (Koen & Eyer 2002; Nazé 2004).

From time-scale arguments, Walborn et al. (2004) suggested that the ‘clock’ underlaying the variability was probably a binary orbit. New light was cast on this issue by Donati et al. (2006a), who found HD 191612 to be only the second O-type star known to possess a magnetic field.\(^2\) Although their observations sampled only a single epoch, they were none the less able to estimate a polar field strength of \sim 1.5 kG (from an observed line-of-sight field of \sim 220 G, by assuming a dipole field), and made the case for 538-d rotational modulation, arguing that the field itself could easily be responsible for the implied slow rotation (through magnetic braking).

Of observational necessity, the magnetic, rotational-modulation model is as yet poorly constrained, and the discussion of XMM–Newton spectroscopy by Nazé et al. (2007) emphasizes a number of discrepancies with the X-ray behaviour expected in the simplest version of this scenario. On the other hand, the alternative orbital model has not been subject to any strong tests (arguably, even a strict, coherent spectroscopic periodicity has yet to be demonstrated robustly). Here, we present the results of an extensive campaign of optical spectroscopy, carried out in an attempt to shed light on these issues.

2 OBSERVATIONS

The major part of our campaign was conducted during the 2004 and 2005 observing seasons. Because of the \sim 18 month variability time-scale, scheduled observations at common-user facilities were generally impractical, and our observations were obtained through service programmes; by taking advantage of telescope time awarded to other scheduled programmes; and by exploiting the goodwill of colleagues. The main data set (166 digital observations spanning 17 yr), summarized in Table A1, is therefore quite heterogeneous.

None of the less, the spectra can be conveniently characterized by wavelength range (‘red’, including Hα 6563 Å; ‘blue’, generally including at least the \sim 4400–4700 Å region) and by resolution (‘high’, \R \gtrsim 4 \times 10^3; ‘intermediate’, \R \gtrsim 4 \times 10^2; and ‘low’). With a few exceptions, the spectra are generally reasonably well exposed, with signal-to-noise ratios (S/Ns) typically approaching \sim 100.

All spectra have been put on a heliocentric velocity scale, and all the H\alpha spectra have been corrected for telluric absorption by division, in topocentric space, with an appropriately scaled and smoothed telluric ‘map’ constructed from high-quality, high-dispersion echelle spectra. (Because the telluric lines are unresolved, direct scaling in optical depth is impossible; we scaled in observed intensity, but checked that scaling in observed, pseudo-optical depth gives negligibly different results.)

3 THE 538-D PERIOD: H\alpha VARIABILITY

As already noted by Nazé et al. (2007), the spectral lines can, for the most part, be separated into two groups according to variability characteristics: the absorption lines of metals and of He II show, at most, small changes in line strength, while the hydrogen and He I lines show large equivalent-width changes. H\alpha shows the largest-amplitude variations of any spectral feature followed in our cam-

\(^2\) The first was \theta^1 Ori C (Donati et al. 2002).

Table 1. H\alpha variability: best-fitting parameters for the arbitrary functional form described by equation (1).

| Parameter | Value | Unit |
|-----------|-------|------|
| \W | 2.51 | ± 0.18 Å |
| \A | 6.74 | ± 0.16 Å |
| \P_{\alpha} | 537.6 | ± 0.4 d |
| \i_0 | JD 245 3415.2 | ± 0.5 |
| \sigma_{\phi} | 0.177 | ± 0.005 |
| \phi_0 | 0.337 | ± 0.005 |

3.1 Ephemeris

Even casual inspection shows systematic variations in H\alpha that repeat on a period close to the \sim 540 d time-scale found in Hipparcos photometry (over only about 2 cycles). In order to quantify that period and its uncertainty, we fit an ad hoc analytical function to the equivalent-width measurements; we find that a truncated Gaussian,

\[
W_\alpha(\phi_0) = W_0 - A \exp\left(\frac{-\phi_0^2}{2\sigma_\phi^2}\right) + \phi_0 < \phi_0 < +\phi_0
\]

where phase zero corresponds to peak H\alpha emission.\(^3\)

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The phase-zero epoch is chosen to be close to the median date of the observations (to minimize the formal error on \i_0), and so, because the density of observations has increased with time, many spectra were obtained at negative epochs. For such observations, we use a logarithmic-like notation for phases, such that \phi_0 = T.23 means phase plus 0.23 in cycle −1.

3.2 H\alpha properties

Fig. 1 shows the H\alpha equivalent-width measurements, which vary between \sim 1.4 and \sim 4.3 Å, folded on the adopted ephemeris; it illustrates a number of noteworthy points.

(i) The H\alpha light curve is remarkably symmetrical about phase zero.

(ii) There is a well-defined interval of apparent ‘quiescence’, lasting \sim 0.3 (\sim 1–2\phi_0) of the period.

(iii) The behaviour is repeatable from the earliest quantitative data (1982),\(^4\) supporting a truly periodic underlying 538-d ‘clock’.

\(^3\) This differs from the Walborn et al. (2004) ephemeris (which was tied to the less well determined minimum in the Hipparcos photometry) by half a cycle.

\(^4\) We measured the equivalent width from a digitized version of the plot of H\alpha given by Peppel (1984), who observed on JD 244 5211, 24 yr before our last observation. We included this point in the fit of equation (1), but its exclusion makes no important changes to the ephemeris.
Figure 1. \( \text{H}\alpha \) measurements folded on the ephemeris of equation (2), shown over two cycles. Upper panel: equivalent-width [the solid line is the ad hoc functional fit, equation (1), described in Section 3.1]. Middle panel: FWHM (upper groups of points) and central velocity (lower) of excess emission. Bottom panel: the Hipparcos photometry (with a scaled, vertically shifted version of the \( \text{H}\alpha \) functional fit to guide the eye).

(iv) None the less, the \( \text{H}\alpha \) profiles are evidently not strictly repeatable; the standard deviation of \( W_\lambda \) measurements about the functional fit is only 0.16 Å, but scrutiny of the spectra shows that real, small-amplitude variability contributes to this dispersion at all sufficiently well-sampled phases, including quiescence, on time-scales longer than a few days. We suspect that any O-type star observed as extensively and intensively as HD 191612 may show \( \text{H}\alpha \) 'jitter' at a similar level (cf., e.g., Kaper et al. 1997; Morel et al. 2004).

The \( \text{H}\alpha \) profile during quiescence is asymmetric and filled in by stellar-wind emission (Fig. 2). It is intermediate in appearance between the profiles of HD 36861 [\( \lambda \) Ori; O8 III(\( \ell \))], pure absorption] and HD 175754 [O8 II(\( \ell \))], P Cygni], although the closest matches among the spectra at our disposal are with HD 193514 [O7 Ib(\( \ell \))] and HD 209975 [O9.5 Ib; Fig. 2].\(^5\) Thus while we have no exact match to the quiescent spectrum, the \( \text{H}\alpha \) profile in that state seems to be largely unremarkable, with no evidence of substantial excess emission relative to normal stars of broadly comparable spectral type.

At other phases, the line-profile morphology is P Cygni like, but the increase in emission is not necessarily associated with an increase in the global mass-loss rate (cf. Section 4.2). Phenomenologically, the changes in the appearance of the profile can be entirely accounted for by variable amounts of roughly Gaussian emission superimposed on a constant, underlying quiescent-state spectrum (Fig. 2). We have characterized the mean velocity displacement and width of this excess emission by Gaussian fits; results are incorporated in Fig. 1. Significant phase dependence is evident for neither mean velocity nor full width at half-maximum (FWHM), which average +7 and 271 km s\(^{-1}\), respectively (standard deviations of 10 and 16 km s\(^{-1}\), commensurate with likely observational uncertainties); the width of the excess emission is much less than both the stellar-wind terminal velocity (Section 6.4), and the ‘windy’ \( \text{H}\alpha \) emission normally seen in luminous O stars.

Finally, although the \( \text{H}\alpha \) and Hipparcos light curves are in phase (to within the errors), we note that the photometric variations are too large to result from excess line emission alone; most of the \( \sim3\)–4 per cent change must arise from true continuum-level variations.

4 THE 538-D PERIOD: OTHER VARIABLE LINES

4.1 Balmer and He\( \ell \) lines

Other Balmer lines (to at least H10)\(^6\) and the He\( \ell \) lines all show large variations, in both strength and velocity, which are reproducible on the 538-d period. He\( \ell \) 4471 Å (Fig. 3) exemplifies the typical behaviour of the He\( \ell \) lines; most remain in absorption at all phases, although He\( \ell \) \( \lambda\lambda \) 5876, 7065, 7281, as well as \( \text{H}\alpha \) and \( \text{H}\beta \), show strong P-Cygni-like emission at maximum. While most lines are in absorption during quiescent phases (even \( \text{H}\alpha \), as illustrated with other Balmer lines in Fig. 2), He\( \ell \) 6678 Å is exceptional in retaining an emission core (Fig. 4).

The basic phenomenology of the variability is, evidently, infilling of a near-normal underlying absorption profile by slightly (\( \sim20–30\) km s\(^{-1}\)) redshifted, almost symmetrical emission (cf. Fig. 2). This interpretation is consistent both with direct inspection of line profiles

\(^5\) We also examined the spectrum of HD 225160; surprisingly, this O8 Ib(\( \ell \)) star shows much stronger emission than the O7 and O9.5 Ib stars.

\(^6\) Paschen lines from 3–8 to 3–20 are recorded in emission in the CFHT ESPaDOnS spectra; these are also strongly variable, in a manner consistent with the 538-d period.
and with the phase dependence of variability: all lines share a common velocity around \( \phi \approx 0.5 \), becoming increasingly blueshifted with decreasing line strength (i.e., increasing emission infill). Moreover, the amplitude of radial velocity variation correlates with the amplitude of line-strength variation.

This phenomenology appears to be applicable to almost all variable features, embracing not only the hydrogen and He I lines (including \( \lambda \lambda 6563 \)), but also the signature of\( \text{O}^+ \text{P} \text{C} \text{III} \lambda \lambda 4650 \) transitions (Fig. 4; the adjacent \( \text{N} \text{II} \lambda \lambda 5001/5005 \) lines are much less, if at all, variable, and so are evidently dominated by ‘normal’, photospheric, Of emission). \( \text{HeII} \lambda 4686 \) shows an exceptional pattern of variability; narrow emission is present even at quiescent phases, and while the overall emission-line strength increases near phase zero, only small changes occur at the position of this narrow emission (Fig. 4), possibly because the emission is already optically thick here.

### 4.2 Ultraviolet P-Cygni profiles

Walborn et al. (2003) presented a high-resolution International Ultraviolet Explorer (IUE) spectrum obtained on 1992 December 19 (JD 244 8975.9, \( \phi = 9.74 \), \( R \approx 10^5 \), \( \lambda \lambda \sim 1200–1900 \) Å). There are no other high-dispersion ultraviolet (UV) spectra available, but the IUE archive contains a low-resolution, short-wavelength spectrum (1984 July 25, JD 244 5907.0, \( \phi = 14.04 \), \( R \approx 300 \)). After smoothing and binning the high-resolution spectrum to render it comparable to the low-resolution one (Fig. 5), there is no evidence for large changes in the resonance-line profiles of \( \text{N} \text{V} \lambda 1240 \), \( \text{SiIV} \lambda 1400 \) and \( \text{CIV} \lambda 1550 \) (although we cannot rule out variations at the moderate level observed in the oblique rotator \( \theta^1 \) Ori C; Walborn & Nichols 1994). This encourages the view that the H\( \alpha \) variability results from the changing visibility of a constrained plasma, rather than a large-scale global change in outflow characteristics – for example, if the variability were a consequence of changes in mass-loss rate, this would be expected to have a clear signature in the UV P-Cygni profiles.

### 5 THE CONSTANT LINES

In contrast to the hydrogen and He I lines, the absorption lines of metals and of \( \text{Heii} \), together with many selective emission lines, show only small changes, at most, in line strength. For simplicity, we label such lines as ‘constant’; if there is any fine-strength variability, it is at a very low level.

#### 5.1 Absorption lines

The \( \text{CIV} \lambda \lambda 5801, 5812 \) Å doublet exemplifies the behaviour of the constant absorption lines. This doublet is particularly well suited to measurement both on astrophysical grounds (the lines are very symmetrical, and are formed deep in the atmosphere, and hence are less likely to be contaminated by ‘windy’ emission than many other transitions (cf., e.g., Fullerton, Gies & Bolton 1996), and observationally (the nearby diffuse interstellar bands at 5778/5780/5797 Å are of similar strength to the \( \text{CIV} \) lines, and provide a very useful zero-point calibration for the wavelength scale, allowing rather precise differential velocities to be obtained even from intermediate-dispersion data of unexceptional quality).

The \( \text{CIV} \) measurements are presented in Fig. 6, and illustrate the typical behaviour of essentially constant line strength coupled with small-amplitude radial-velocity variations. The steady increase in velocity between JDs \( \sim 2453500 \) and 245 3700 discussed by Nazé et al. (2007) clearly does not repeat on the \( 538 \)-d period. Metal absorption lines strong enough to be measured consistently in most spectra (e.g. \( \text{NII} \lambda \lambda 4515–4554 \), as well as the \( \text{HeII} \) absorption lines (e.g. \( \lambda \lambda 4200, 4541, 5411 \)), follow essentially the same behaviour as the archetypal \( \text{CIV} \) lines, as illustrated in Fig. 3 by data for \( \text{HeII} \lambda 4541 \).

#### 5.2 Selective emission lines

The high-quality UES and CFHT spectra, in particular, reveal a rich spectrum of selective emission lines (cf. Walborn 2001). In general, the weakness of the lines precludes detailed scrutiny of their phase dependence in the remaining data, but comparison of quiescent and emission-line phases in these echellograms establishes that most of the emission features’ exhibit essentially the same behaviour as the...
C IV lines (i.e., near-constant line strength and small-amplitude radial velocity variations). Only the stronger lines can be consistently measured, but results for Si IV 6667 Å shown in Fig. 6 illustrate the general accord between the behaviours of the selective emissions and the ‘constant’ absorption lines.

The extensive far-red coverage of the high-quality CFHT ESPaDOnS spectra allows many other weak emission lines to be recognized; these appear also to show little or no variability in line strength (at least, between the two epochs sampled, at 538-d phases 0.24 and 0.89; the ‘variable’ lines discussed in Section 4 change substantially between these epochs). Most of these features are previously unreported in O star spectra, and currently lack persuasive identifications. Measured wavelengths in Å [and possible identifications, with multiplet numbers from Moore (1945)] are, for the stronger features, 5739.8, 6394.8, 6467.1, 6478.7 (N III 14? good wavelength matches but not all multiplet members present); 6482.3;
phase in the H\textsc{iue} Figure 5.

Although $\lambda$ Figure 4.

Emission-line variability in HD 191612; spectra are labelled by $\phi\alpha$ spectrum, $\phi\alpha$ complex; the unlabelled tick marks indicate wavelengths of O\textsc{ii} lines. Note data, $\lambda\lambda$ amplitude. The quiescent-state spectrum) is notable only for the unusually large relative (i.e. growth of a slightly redshifted, fairly narrow emission superimposed on emission throughout ‘quiescence’, the qualitative nature of the variability not to be consistent with a standard C\textsc{iii} identification.

$\phi\alpha = 6728.8 \pm 0.04$; upper spectrum (offset by 0.5 continuum units): high-resolution spectra of HD 191612. Lower spectrum: low-resolution, $\phi\alpha = 9.74$, smoothed and binned to match the low-resolution data.

5.3 Radial-velocity variations and spectroscopic orbit

The ‘constant’ lines show significant, systematic radial-velocity variations$^8$ which do not repeat with the 538-d period (Figs 3 and 6; Table A2). These velocity variations cannot reflect photospheric motion about a static centre of mass (the implied change in radius is as great as $\sim 250 R_\odot$, or $\sim 17 R_\odot$). In principle, they could be attributed to changes in the depth of line formation in an accelerating outflow, through density changes in or near the transonic region resulting from stochastic changes in the stellar-wind mass-loss rate. However, the lack of changes in H$\alpha$ in mean quiescent spectra around phases $\tilde{1}$.5 and 0.5 (corresponding to negative- and positive-velocity C\textsc{iv} states) does not encourage confidence in this interpretation, and we have already argued that the C\textsc{iv} lines are formed relatively deep in the subsonic atmosphere.

We therefore consider the possibility of orbital motion on periods other than that of the major spectroscopic variability. We concentrate on measurements of the C\textsc{iv} $5800$ Å doublet and the $\sim 6700$ Å emission-line complex (Si\textsc{iv} + unidentified), which yield mutually consistent results. Both features have velocity zero-points tied in to nearby, moderately strong Diffuse Interstellar Bands (we adopted DIB wavelengths as measured in the CFHT spectra), and therefore record the velocity variations more precisely and accurately than other lines that lack this advantage. Results are given in Table A2.

Fortuitously, the data set includes good-quality observations showing that relatively large positive velocities occurred in 538-d cycles $-3$ and $-6$, as well as at $\phi\alpha \sim 0.5$, hinting at $P_{\text{orb}} \simeq 3P_{\text{c}}$. Trial orbital solutions with $P_{\text{orb}} \simeq 1500–1600$ d yielded residuals that are satisfactorily small, but none the less slightly larger than the formal errors returned by the Gaussian fits to lines used to measure velocities. Since external errors evidently dominate (C\textsc{iv}) or match (Si\textsc{iv}) the formal uncertainties, we weighted all C\textsc{iv} measurements and all Si\textsc{iv} measurements equally, but with the Si\textsc{iv} measurements assigned approximately one-third weight to reflect their greater residuals (and adjusted by $-2.1$ km s$^{-1}$ to bring them to the same $\gamma$ velocity as the C\textsc{iv} measurements).

The resulting orbital solution is summarized in Table 2, and illustrated in Fig. 7. While this solution has not yet been subject to the acid test of predictive power, its success in reproducing observations at three separate periastron passages, with small residuals that are consistent with realistic observational uncertainties, encourages the view that the solution does characterize a true binary orbit.

5.3.1 Secondary orbit

A careful examination of weak lines of low-ionization stages in the best data shows velocity displacements in antiphase with the stronger lines, bolstering the interpretation of the 1540-d velocity variations as being orbital in origin. The spectroscopic fingerprint of the secondary is weak, but its velocity can be reasonably well quantified by simultaneous Gaussian fits to a selection of a half-dozen unblended, relatively strong O\textsc{ii} absorption lines ($\lambda\lambda 4300–4700$ Å, central depths 2–3 per cent below continuum),

$^8$ The narrow core of He\textsc{i} $\lambda 4686$ emission, present throughout quiescence, tracks the motion of the ‘constant’ absorption lines and selective emission lines.

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Figure 6. Equivalent-width and velocity variations for the C\textsc{iv} 5801 Å absorption and Si\textsc{iv} 6667 Å emission lines; other details are as for Fig. 3.

Although we do not have the benefit of well-defined interstellar features to provide a velocity reference for most observations encompassing the O\textsc{ii}-line spectral region, we can exploit the fact that the centrally placed He\textsc{ii} 4541 Å line follows the motion of the C\textsc{iv} and Si\textsc{iv} features used to determine the primary’s orbit (Section 5.1). By measuring O\textsc{ii} velocity differences with respect to λ4541, and correcting for computed motion of the primary, we can in effect remove spectrum-to-spectrum velocity zero-point offsets. (Using directly observed velocities gives essentially identical results, but with somewhat larger scatter.)

An orbital solution of the resulting O\textsc{ii} velocities gives γ and $K_2$ (with all other parameters fixed at the primary-orbit values$^9$); the

$^9$For completeness, we report that γ velocity for the O\textsc{ii} lines is, formally, +9.1 ± 1.0 km s$^{-1}$. Although small differences in γ velocities for different lines would not be unusual in an O-type star, it should be noted that our γ velocities for the C\textsc{iv} doublet and the Si\textsc{iv}/6700 complex depend on measured (not laboratory) wavelengths for both the reference DIBs and the unidentified λ6700 emissions, while the weakness of the O\textsc{ii} lines means that small zero-point shifts arising from, for example, unrecognized blends would not be surprising.
slope of He II versus He II–O II velocities gives an entirely consistent mass ratio. Results are incorporated into Table 2.

5.3.2 Relationship between the 538- and 1540-d periods?

The orbital period is $\sim 3 \times 538$ d, which has suggested to colleagues the possibility of some sort of resonance. However, a solution with $P_{\text{orb}} = 3 P_u$ gives a significantly poorer fit than one with $P_{\text{orb}}$ allowed as a free parameter.\(^{10}\) Thus, although it seems to be quite securely established that HD 191612 is a long-period binary, it appears that, with $P_{\text{orb}} = 2.87 \pm 0.03 P_u$, the binary orbit has no direct important role in the large-amplitude spectroscopic variability.

There is, however, an interplay between the periods in a limited observational sense, explaining otherwise complex behaviour seen in some lines. Fig. 8 shows spectra obtained near the extremes of both the orbital motion and the spectroscopic variability. In addition to illustrating the small orbital velocity amplitudes (less than the linewidths) the quiescent-state spectra clearly indicate that absorption in the Si II $\lambda\lambda4552, 4568, 4575$ triplet tracks the secondary spectrum, on the orbital period. The primary spectrum has emission in these lines which varies on the 538-d period, and which shows a strong decline in strength going from $4.552$ to $4.575$ (as is observed in some LBV/WN11 spectra; e.g. He 3-519, Walborn & Fitzpatrick 2000).

The interplay of these two cycles means that the lines can appear purely in absorption (throughout quiescent 538-d phases), but in other states the behaviour is more complex. When the primary and secondary spectra are at or near maximum velocity separation, a P-Cygni-like appearance results in $\lambda\lambda4552, 4568, 4575$ with absorption in $\lambda4575$ (as in the first [top] spectrum in Fig. 8); emission phases at smaller velocity separations can result in apparent disappearance, or nulling, of $\lambda4552$ (third spectrum in Fig. 8).

A similar effect is detectable in the N II spectrum, in particular the 5001/5005 Å lines, where primary emission and secondary absorption clearly move in antiphase; and, possibly, in the $\lambda\lambda4414.9, 4417.0$ O II lines.

5.4 System characteristics

Mass constraints implied by the orbital solution are illustrated in Fig. 9. We have no constraint on the orbital inclination, but the masses are entirely consistent with a system comprising a $\sim 30 M_\odot$, late-O giant accompanied by a $\sim 15 M_\odot$, early-B main-sequence star, viewed at an intermediate angle ($\sin i \approx 0.5$). The projected centres-of-mass separation at periastron is $(1 - e) a \sin i = (555 \pm 68) R_\odot \sim 40 R_\odot$; although this certainly allows for the possibility of wind–wind interactions, the dominant H$\alpha$ emission region, for example, is surely much closer to the primary, so that it is unlikely that the secondary plays any important role in the 538-d changes.

The strong variability of the He I classification lines, together with the small orbital-velocity amplitudes (which ensure that the

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\(^{10}\) An $F$-test yields a probability of $< 1$ per cent that $\chi^2$ for the fixed-period fit is no poorer than that for the solution with $P_{\text{orb}}$ free; that is, formally rules out that $P_{\text{orb}} = 3P_u$ with $> 99$ per cent confidence.

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**Table 2.** Orbital solution. The main orbital parameters are constrained by measurements of C IV 5801 Å and Si IV 6667 Å in the primary spectrum; $K_2$ is established from O II lines in the secondary spectrum.

| Parameter | Value |
|-----------|-------|
| $v$       | $-5.19 \pm 0.36$ km s$^{-1}$ |
| $K_1$     | $11.77 \pm 0.84$ km s$^{-1}$ |
| $e$       | $0.438 \pm 0.038$ |
| $\omega$  | $344.7 \pm 6.5$ |
| $P_{\text{orb}}$ | $1542 \pm 14$ d |
| $T_0$     | JD 245 3720 |
| $f(m)$    | $0.190 \pm 0.042 M_\odot$ |
| $a_1 \sin i$ | $322 \pm 24 R_\odot$ |
| rms residual (weight 1, C IV) | $2.2$ km s$^{-1}$ |
| $K_2$     | $24.4 \pm 1.4$ km s$^{-1}$ |
| $f(m)$    | $1.68 \pm 0.29 M_\odot$ |
| $a_2 \sin i$ | $667 \pm 38 R_\odot$ |
| $q = M_2/M_1$ | $0.483 \pm 0.044$ |
| rms residual (O II) | $5.1$ km s$^{-1}$ |

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**Figure 8.** Signatures of the secondary’s spectrum. Data have been lightly smoothed, rebinned, and renormalized to facilitate comparison between observations, which are labelled by orbital/H$\alpha$ phases (Section 5.3/3.1). Spectra have been shifted to bring the primary to zero velocity; the long tick marks indicate rest wavelengths of the identified lines. In this reference frame, the He II, N II and Si IV lines show no significant movement (i.e., arise principally in the primary’s spectrum). The short tick marks show the expected positions of the O II and Si IV lines in the secondary spectrum, according to the orbital solution given in Table 2.
stars are sufficiently spectroscopically peculiar that any luminosity-class argument based on direct comparisons with normal stars must be considered hazardous.

6.2 Spectroscopic fine analysis

Although the peculiarities of the spectra preclude extremely detailed modelling, the basic photospheric parameters can none the less be quite well constrained. We searched for matches to the ‘quiescent’ spectrum of 2001 August 5 ($\phi_0 = 3.61$), in a grid of FASTWIND models (Puls et al. 2005) sampled at steps of 1 kK and 0.2 in $T_{\text{eff}}$ and log $g$ (with $Y = 0.1$ and $v_{\text{turb}} = 15$ km s$^{-1}$). We find a generally good match with observations for a model with $T_{\text{eff}} = 35.0$ kK, log $g = 3.5$, in excellent agreement with the analysis reported by Walborn et al. (2003, based on simpler models and poorer-quality data).

No model reproduces the strength of HeI 4471 Å (cf. the general discussion of this issue by Repolust, Puls & Herrero 2004), but the HeI singlet lines are fully consistent with the estimated $T_{\text{eff}}$; adopting 34 kK would still require log $g = 3.5$, but leads to the model HeII lines becoming too weak. The best-fitting model parameters can therefore be considered as fairly well determined, to $\pm 1$ kK, 0.1 in $T_{\text{eff}}$, log $g$ (although, of course, no model matches the peculiar HeII 4686 and HeI 6678 Å profiles).

This takes no account of the effect of the secondary spectrum, which is expected to have somewhat stronger HeII lines, while the HeI line strengths may be underestimated by $\sim 10$ per cent through dilution. This would result in the primary being, if anything, modestly warmer (by perhaps $\sim 1$ kK) than the foregoing analysis suggests.

6.3 Reddening, radius

To determine the reddening and angular diameter, we compared a $T_{\text{eff}} = 35.0$ kK, log $g = 3.5$ OSTAR2002 model (Lanz & Hubeny 2003) with archival low-resolution IUE spectrophotometry and optical & 2MASS photometry (Flux variability is of too small an amplitude to be of concern here). We estimate $(R_d/R_\odot)/(D/kpc) = 6.7 \pm 0.05$ and $E(B-V) = 0.56 \pm 0.03$ [using a Cardelli, Clayton & Mathis 1989 reddening law with $R_V \equiv A_V/E(B-V) = 3.1$], where $R_d$ is an ‘effective’ radius characterizing the emitting surfaces. The match between the model and observations, though not perfect, is reasonably good, and there is no evidence of any infrared excess out to 2 μm. The fit is slightly improved by allowing for the contribution of a cooler secondary, such as a $\sim 20$ kK secondary contributing $\sim 10$ per cent of the light at $V$, and this ad hoc model is illustrated in Fig. 10.

The reddening and general properties are consistent with the membership of the Cyg OB3 association, as noted by Humphreys (1978); we adopt her association distance of 2.29 kpc [see discussion in Walborn (2002) for the embedded cluster NGC 6871], whence the primary’s radius is

$$R_\star \simeq 14.5 \left( \frac{f_1}{0.9} \right) \left( \frac{d}{2.3 \text{kpc}} \right)^{35 \text{kK}} \frac{10^{0.2(E(B-V)/1.0)}}{R_\odot},$$

where $f_1$ is the fractional V-band luminosity of the primary and the $T_{\text{eff}}$ term accounts for the approximate $T_{\text{eff}}^2$ scaling of model-atmosphere V-band surface fluxes in this temperature range. The primary then has log $(L/L_\odot) \approx 5.4$ and $M(V) \simeq -5.6$ – parameters broadly consistent with a late-O giant.
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Figure 10. Spectral energy distribution for HD 191612 (red). The weighted sum of an OSTAR2002 model with $T_{\text{eff}} = 35$ kK, $\log g = 3.5$ and an ATLAS9 model with $T_{\text{eff}} = 20$ kK, $\log g = 4.5$, with a $V$-band flux ratio of 9:1 and $E(B-V) = 0.56$, is shown for comparison (black; see Section 6.3 for details).

6.4 Hα mass-loss rate

In our FASTWIND models, the mass-loss rate is characterized by a parameter

$$ Q = \frac{M \sqrt{f_{cl}}}{\text{M}_\odot \text{yr}^{-1}} \left( \frac{v_{\infty}}{\text{km s}^{-1}} \right)^{-1.5} $$

(Puls et al. 2005), where $f_{cl} = \langle \rho^2 \rangle / \langle \rho \rangle^2$ is the clumping factor; from modelling of Hα, we find $-\log Q = 12.6-12.7$ for the minimum-state spectrum. We estimate $v_{\infty} \approx 2700$ km s$^{-1}$ (with $\sim 10-15$ per cent uncertainty) from the C IV resonance doublet in the IUE spectra, whence the primary’s quiescent stellar-wind mass-loss rate is $\log(M \sqrt{f_{cl}}) \approx -5.8$ dex $\text{M}_\odot \text{yr}^{-1}$. As for other aspects of the FASTWIND analysis, this is in satisfactory agreement with the corresponding result reported by Walborn et al. (2003; $M = -5.6$ dex $\text{M}_\odot \text{yr}^{-1}$).

6.5 Rotation

The equatorial velocity expected for a 538-d rotation period and $R_\ast \approx 14.5 R_\odot$ is $\sim 1.4$ km s$^{-1}$, and a measurement of $v_c \sin i$ would provide a strong test of rotational modulation. (For reference, the equatorial rotation velocity for synchronous rotation is $\sim 0.5$ km s$^{-1}$, or, for pseudo-synchronous [periastron-synchronized] rotation, $\sim 0.8$ km s$^{-1}$.) Most lines in the spectrum show asymmetries or variability to some extent but, as already discussed, the C IV 5800 Å doublet shows exceptionally symmetrical profiles with little or no evidence of contamination by emission; these high-excitation lines are expected to form rather deep in the atmosphere, and therefore to be more representative of the subsonic photosphere than many other features.

Donati et al. (2006a; see also Howarth 2003) showed that the C IV profiles deviate strongly from those expected from rotational broadening, but are well matched by simple Gaussian ‘turbulence’, with zero rotation. We have attempted to set upper limits to the rotation rate by looking for zero-amplitude nodes in the Fourier transform of the profiles (Gray 1992) in the combined 2005 CFHT spectra ($S/N > 10^3$). We compared the results with transforms of synthetic spectra that include rotational broadening, isotropic Gaussian macroturbulence, and Gaussian noise, using an intrinsic profile from an OSTAR2002 model (Lanz & Hubeny 2003).

Illustrative results are shown in Fig. 11. Even for these very high-quality results there is no sign of rotational zeroes above the noise...
(which dominates at frequencies \( \gtrsim 10^{-2} \, \text{s km}^{-1} \)). Empirically, we find that the observations are well matched by any model having

\[
\sqrt{\sigma_v^2 + \left( \frac{v_e \sin i}{2} \right)^2} \simeq 45 \, \text{km s}^{-1}
\]

(where \( \sigma_v \) characterizes the velocity dispersion of the macroturbulence and \( v_e \sin i \) characterizes the projected equatorial rotation velocity), for \( v_e \sin i \lesssim 60 \, \text{km s}^{-1} \); larger values of \( v_e \sin i \) are ruled out by the data. This is a disappointingly weak upper limit, given the data quality, but results from the effectiveness of the Gaussian turbulence in removing information with high spatial frequencies from the profiles.

7 DISCUSSION

The symmetry and reproducibility of, in particular, the H\( \alpha \) 538-d light curve strongly suggest an origin in an essentially geometrical process – that is, the changing aspect of a distinct emission region. This is also consistent with the absence of evidence of associated global changes in stellar-wind properties (Section 4.2). Since we have found that the variability is not orbital, our results lend strong support to the proposal by Donati et al. (2006a) of rotationally modulated line emission from a magnetically constrained plasma.

One might hope that the fairly rich emission spectrum would offer insight into physical conditions in the line-forming region, but there is a dearth of traditional nebular diagnostics (suggesting a high-density regime). We can only infer some general characteristics of the H\( \alpha \)-emitting region. First, optically thin H\( \alpha \) emission in any time-independent axisymmetric structure cannot account for the observed 538-d behaviour; at least half the emission from any such structure would be visible at all times, but emission equivalent to the adopted ‘excess’ emission cannot plausibly be present during ‘quiescent’ phases. At the same time, very complex geometries (such as observed in \( \tau \) Sco; Donati et al. 2006b) are unlikely, as they would not be consistent with the large amplitude of emission variability.

Secondly, for the distance and reddening adopted in Section 6.3, and assuming isotropic, case-B recombination emission, the maximum excess H\( \alpha \) emission over the quiescent state corresponds to

\[
\int n_e^2 \, dV \simeq 7 \times 10^{22} \, \text{cm}^{-6} \, R_\odot^2.
\]

This is a lower limit to the true emission measure if the emission is not optically thin.

If we adopt the hypothesis that the geometry of the line-emitting region is determined by the magnetic field, then the simplest acceptable magnetic field geometry is a misaligned, centred dipole – known to provide a reasonable approximation to \( \theta^1 \) Ori C, the prototype magnetic O star (Stahl et al. 1996; Donati et al. 2002; Wade et al. 2006). Babel & Montmerle (1997) argue that, for such a geometry, stellar-wind material from the magnetic polar and temperate regions will be deflected along the field lines towards the magnetic equator, where the colliding flows shock to create a high-temperature, X-ray-emitting plasma. The plasma cools to form a dense disc which gives rise to optical emission. Magnetohydrodynamic calculations support this general outline, albeit with a number of refinements (e.g. ud-Doula & Owocki 2002; Gagné et al. 2005; Ud-doula, Townsend & Owocki 2006).

Since this model is at least consistent with the direct magnetic field measurements of HD 191612 (Donati et al. 2006a), for heuristic purposes we explore schematic ‘toy’ models of H\( \alpha \) emission from a centred, tilted, geometrically thin disc, in which the relative emission is simply related to the projected area (taking into account limb darkening and occultation by the star). Although this is a physics-free, geometrical model, we can speculate that it may characterize a more realistic scenario; for example, for a disc-like emission region, the ‘limb darkening’ may actually correspond to increasing H\( \alpha \) optical depth as the line of sight approaches the plane of the disc.

We find that this simple model can provide a reasonable match to the H\( \alpha \) variability, provided that (i) the sum of the inclination of the rotational axis to the line of sight and the angle between that axis and the magnetic axis is close to 90°, and (ii) moderately strong limb-darkening is present (so that the emission is low at all rotational phases when the line of sight is close to the disc plane). Moreover, the width-to-thickness ratio must be reasonably large, to account for the large amplitude of emission-line variability.

A specific model, selected through a genetic-algorithm minimization, is shown in Fig. 12; its parameters are inner radius \( R_{\text{in}} = 2.00 \, R_\odot \), outer radius \( R_{\text{out}} = 2.13 \, R_\odot \), axial inclination \( i = 68° \), magnetic-axis offset \( \alpha = 27° \), and linear limb-darkening coefficient \( u = 0.7 \). These parameters should be regarded only as illustrative, as many other combinations give closely similar fits (e.g. 1.55\( R_\odot \), 1.76\( R_\odot \), 41°, 56°, 0.6 gives results almost indistinguishable from those in Fig. 12). The radii, in particular, are only very weakly constrained from the H\( \alpha \) light curve, because the total emission is fixed by an arbitrary scaling factor (parametrizing the surface emissivity).

[In principle, the continuum photometry could help fix the radii, but in practice it only weakly constrains the model, because the photometric amplitude is so small, the noise is relatively high, and the number of free parameters is large. The only firm conclusion is that a very large disc, very close to the star, and optically thick in the continuum, is not allowed. The broad-band photometry suggests that any disc cannot contribute more than \( \sim 3–4 \) per cent of the visible-region broad-band (\( H_\beta \)) flux – that is, the implied equivalent width of the H\( \alpha \) emission referred to disc continuum is as much as \( \sim 150 \, \text{A} \).]

The toy model succeeds in reproducing the H\( \alpha \) light curve, but it would clearly be rash to put too much weight on this success. Not only does the physical model which inspired it encounter difficulties in accounting for the details of the observed X-ray emission (Nazé et al. 2007), but also the disc model is certainly not unique. To illustrate this, we have also considered a minimalist model of H\( \alpha \) emission from a single surface spot, described by

\[
f_\alpha = f_0 + A (\cos i \cos \beta + \sin i \sin \beta \cos \phi)(1 - u + u \cos \mu),\]

11 Though it lacks any compelling physical basis, this ‘spot’ model was partly motivated by analogy with magnetic oblique rotators; of course, a centred oblique dipole would be expected to give rise to two ‘spots’, one of which would be visible at any phase.
where $f_0$ and $A$ are normalizing constants, $i$ is the inclination of the rotation axis to the line of sight, $\beta$ is the colatitude of the spot, $\phi$ is the rotational phase, $u$ is a linear limb-darkening coefficient, and $\mu$ is the angle between the line of sight and the surface normal at the spot. We find that the form of the Hα light curve can again be well matched by this model, with $u \simeq 0.6$, $i + \beta \simeq 105^\circ$, $i \simeq \beta$ (Fig. 12).

8 SUMMARY AND CONCLUSIONS

We have shown that the O6.5f?pe–O8fp spectroscopic variations observed in HD 191612 are underpinned by an extremely regular 538-d ‘clock’ which has kept good time across 24 yr of quantitative data (and which is documented a decade farther back by photographic material). The Balmer and He I lines show changes which are highly reproducible on this period, and which are characterized by variations in slightly redshifted, moderately narrow ‘excess’ emission. Absorption lines of metals and He II are essentially constant in line strength (as are many selective emission lines), but show radial-velocity changes arising in a double-lined spectroscopic-binary orbit with $P_{\text{orb}} = 1542$ d. The components’ properties are broadly in accord with an O8 giant-like primary and a B1 main-sequence companion.

The results are entirely consistent with the 538-d variability arising through rotational modulation of a magnetically constrained plasma, as proposed by Donati et al. (2006). As in the case of the B0.2 V star τ Sco (Donati et al. 2006b), the implied slow rotation and the long-term stability of the variations argue that the field originates in a fossil remnant, rather than a dynamo. However, although toy ‘disc’ models (among others) are consistent with some aspects of the data, we are unable to constrain the geometry of the emission region in an interesting way. Future work could therefore usefully concentrate on improving our understanding of the field geometry, which should provide a better framework for numerical models of optical emission, and perhaps help resolve the issues of anomalously broad X-ray lines and soft X-ray emission reported by Nazé et al. (2007). A continuing programme of circular spectropolarimetry is working towards this aim, and a magnetic-field measurement from the second-epoch ESPaDOnS observation used here already gives a clearly detected field with a longitudinal component larger than the discovery measurement, as expected for the geometry adopted by Donati et al. (2006) and in our toy disc model.

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APPENDIX A: OBSERVATIONAL DETAILS

This appendix summarizes the following.

(i) The log of observations (Table A1, available in full online).
(ii) Radial-velocity measurements used in the orbital solutions (Table A2).
Table A1. A sample of Table A1 (available in full online), the log of optical spectroscopy used in this paper. Each ‘observation’ is a data set collected at a single site in a given night, and may consist of several separate spectra. Phases are calculated with the ephemerides of Table 2 and equation (2); the notation $\tilde{C}_{.37}$ means phase $+.37$ in cycle $-C$. Where two observers are identified, the second name is the PI or instigator of the observation; ‘srv’ indicates a service-mode observation. Where observers provided smoothed or binned spectra, the effective resolution is tabulated ($\Delta\lambda$).

| JD     | Year | Phase $\phi_H$ | Phase $\phi_\pi$ | $\lambda$ range | Telescope/instrument | Observer | $\Delta\lambda$ | $W_1$ (Hz) |
|--------|------|----------------|------------------|------------------|----------------------|----------|-----------------|------------|
| 2447691.5 | 1989.45 | T.T. 34.1 | J.109 | 651–662 | INT/IDS | Prinja/Howarth | 0.5 | +1.37 |
| 2447724.7 | 1989.54 | T.T. 44.11 | J.11 | 404–499 | INT/IDS | Herrero | 0.8 |
| 2447276.5 | 1989.55 | T.T. 44.11 | J.11 | 634–677 | INT/IDS | Herrero | 1.5 | +1.41 |
| 2448217.4 | 1990.62 | T.T. 34.17 | J.37 | 394–474 | INT/IDS | Howarth | 1.5 |
| 2449139.7 | 1993.41 | T.T. 34.03 | J.303 | 439–481 | INT/IDS | Vilchez | 0.5 |
| 2453955.4 | 2006.60 | T.T. 34.15 | J.16 | 637–673 | Tartu | Kolka | 1.5 | -4.00 |
| 2453956.5 | 2006.60 | T.T. 34.15 | J.16 | 637–673 | Tartu | Kolka | 1.5 | -3.96 |
| 2453958.4 | 2006.61 | T.T. 34.15 | J.16 | 637–673 | Tartu | Kolka | 1.5 | -4.14 |
| 2453966.6 | 2006.63 | T.T. 34.16 | J.38 | 384–706 | WHT/ISIS | Leisy/Lennon | 0.4 | -4.07 |
| 2454002.5 | 2006.73 | T.T. 34.18 | J.38 | 384–706 | WHT/ISIS | Leisy/Lennon | 0.4 | -3.32 |

Table A2. Velocities used in orbital solutions. Note that each of the three data sets is subject to a separate velocity zero-point error of the order of a few km s$^{-1}$ (see footnote 9).

| JD | $V_\odot$ (km s$^{-1}$) | JD | $V_\odot$ (km s$^{-1}$) | JD | $V_\odot$ (km s$^{-1}$) | JD | $V_\odot$ (km s$^{-1}$) |
|----|------------------|----|------------------|----|------------------|----|------------------|
| 2449529.4 | -9.3 | 2453203.5 | -11.3 | 2453275.2 | -11.3 | 2453474.6 | -8.8 | 2453601.4 | +4.4 |
| 2452127.5 | +9.0 | 2453225.8 | -6.3 | 2453276.2 | -10.6 | 2453488.6 | -7.3 | 2453625.4 | +3.7 |
| 2452288 | -8.3 | 2453246.3 | -8.9 | 2453278.3 | -11.7 | 2453519.6 | -5.9 | 2453652.3 | +7.0 |
| 2452549.4 | -3.1 | 2453251.3 | -9.6 | 2453282.4 | -12.0 | 2453536.7 | -5.8 | 2453653.3 | +7.6 |
| 2452607.3 | -5.3 | 2453264.3 | -7.8 | 2453283.3 | -10.1 | 2453545.0 | -3.2 | 2453668.3 | +6.3 |
| 2452801.6 | -11.4 | 2453283.3 | -9.6 | 2453302.3 | -11.3 | 2453456.0 | -3.7 | 2453681.3 | +11.7 |
| 2452808.5 | -11.2 | 2453266.3 | -12.2 | 2453314.6 | -10.3 | 2453547.0 | -3.5 | 2453836.8 | +6.3 |
| 2452816.6 | -11.0 | 2453267.2 | -9.9 | 2453316.3 | -9.7 | 2453548.0 | -4.3 | 2453892.7 | -0.2 |
| 2452871.2 | -12.1 | 2453267.4 | -8.2 | 2453320.3 | -12.9 | 2453554.6 | -7.4 | 2453896.0 | +2.1 |
| 2453132.9 | -6.5 | 2453270.2 | -10.9 | 2453324.3 | -12.7 | 2453568.5 | -0.8 | 2453945.5 | -3.4 |
| 2453146.7 | -10.4 | 2453270.2 | -14.2 | 2453326.3 | -13.2 | 2453585.5 | -2.2 | 2453966.6 | -1.7 |
| 2453192.5 | -7.6 | 2453272.4 | -8.3 | 2453346.4 | -9.5 | 2453594.5 | -4.5 | 2454002.5 | -3.2 |

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SUPPLEMENTARY MATERIAL

The following supplementary material is available for this article online.

Table A1. The log of optical spectroscopy used in this paper.

This material is available as part of the online article from: http://www.blackwell-synergy.com/doi/abs/10.1111/j.1365-2966.2007.12178.x

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