Structure and evolution of the first CoRoT exoplanets: probing the brown dwarf/planet overlapping mass regime

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Received 20 February 2009 / Accepted 4 July 2009

ABSTRACT

We present detailed structure and evolution calculations for the first transiting extrasolar planets discovered by the space-based CoRoT mission. Comparisons between theoretical and observed radii provide information on the internal composition of the CoRoT objects. We distinguish three different categories of planets emerging from these discoveries and from previous ground-based surveys: (i) planets explained by standard planetary models including irradiation; (ii) abnormally bloated planets; and (iii) massive objects belonging to the overlapping mass regime between planets and brown dwarfs. For the second category, we show that tidal heating can explain the relevant CoRoT objects, providing non-zero eccentricities. We stress that the usual assumption of a quick circularization of the orbit by tides, as usually done in transit light curve analysis, is not justified a priori, as suggested recently by Levrard et al. (2009), and that eccentricity analysis should be carefully redone for some observations. Finally, special attention is devoted to CoRoT-3b and to the identification of its very nature: giant planet or brown dwarf? The radius determination of this object confirms the theoretical mass-radius predictions for gaseous bodies in the substellar regime but, given the present observational uncertainties, does not allow an unambiguous identification of its very nature. This opens the avenue, however, to an observational identification of these two distinct astrophysical populations, brown dwarfs and giant planets, in their overlapping mass range, as done for the case of the 8 Jupiter-mass object Hat-P-2b. According to the presently published error bars for the radius determination and to our present theoretical description of planet populations and evolution, the high mean density of this object requires a substantial metal enrichment of the interior and is inconsistent at about the 2-sigma limit with the expected radius of a solar-metallicity brown dwarf. Within the aforementioned observational and theoretical determinations, this allows a clear identification of its planetary nature, suggesting that planets may form up to at least 8 Jupiter masses.

Key words. stars: low-mass, brown dwarfs – stars: planetary systems

1. Introduction

The first space-based project devoted to the search for transiting planets, CoRoT, is now collecting its first results. Seven transiting planets have been announced, with the most recent one, CoRoT-7b, being a super-Earth of \( \sim 5 \)–\(10\) \( M_\oplus \), still to be confirmed by radial velocity follow-up. In the present work, presented during the CoRoT international symposium taking place in Paris, we analyse the first confirmed transiting planets of CoRoT. They highlight the existence of three different categories of transiting exoplanets, which have already emerged from the previous ground-based surveys. Such a distinction between three types of objects is simply used for illustration and does not correspond to different genuine populations and evolution. The first category includes planets whose radii can be explained by standard models and for which a bulk composition can be inferred with little ambiguity. CoRoT-4b is a template of this class and is discussed in Sect. 3. The second kind of objects concerns inflated planets with abnormally large radii, which require extra-mechanisms to explain this behaviour. CoRoT-1b and -2b belong to this category and are analysed in Sect. 4. Finally, the last category of objects, illustrated by CoRoT-3b, is populated by massive objects (\( \gtrsim 10 \) \( M_J \)) which lie in the overlapping mass regime between brown dwarfs and giant planets. Special attention is paid to these objects in Sect. 5. We show how their radius determination can provide valuable information on their true nature and thus on their formation mechanism: gravitational collapse of a molecular cloud for brown dwarfs – yielding objects with the same metallicity as the central star – against core accretion in a protoplanetary disk for planets (Pollack et al. 1996; Alibert et al. 2009), the most widely accepted scenario for planet formation, yielding a substantially heavy material enriched object, as for our Solar system giant planets. The alternative so-called gravitational instability scenario for planet formation (Boss 1997) has been shown to be excluded, both on theoretical (Rafikov 2005, 2007) and numerical (Boley et al. 2007, see Dullemond et al. 2009 for a recent review) grounds, at least for planets within about 50 to 100 AUs. In Sect. 2, we first briefly summarize the planet structure and evolution models used to describe the properties of these planets.

2. Model description

The main physics inputs used in the present models have been described in previous works devoted to the evolution of brown dwarfs and giant planets (Baraffe et al. 2003, 2008). Evolutionary models are based on a consistent treatment between non-grey irradiated atmospheric structures and inner
structures. The interior is composed of a gaseous envelope of solar mixture described by the Saumon-Chabrier equation of state (EOS, Saumon et al. 1995). The inner structure models also take into account enrichment of heavy material, with the appropriate EOS (Saumon et al. 1995). The interior is composed of a gaseous envelope of solar composition. In this case, the predicted radius is too low compared to the computation of an isolated gaseous sphere with solar composition (solid line). In this case, the predicted radius is too low compared to the computation of an isolated gaseous sphere with solar composition (solid line). For all the CoRoT objects, $e = 0$ is assumed in the analysis of the light-curve and of the radial velocity data.

3. Standard planets

In this section we focus on planets whose radius can be explained by the “standard” models discussed in Sect. 2. CoRoT-4b belongs to this population. The planetary and stellar properties of Table 1. Transiting Exoplanets characteristics.

| Object         | $M(M_J)$  | $R(R_J)$ | Period (a.u.) | $\log(F_{\text{bol}})$ | $e$ | Spec. Type | $T_{\text{eff}}(K)$ | Age (Gyr) | $M(M_\odot)$ | $[\text{Fe/H}]$ | $R(R_\odot)$ |
|----------------|-----------|----------|---------------|-------------------------|-----|------------|---------------------|-----------|-------------|----------------|--------------|
| CoRoT-1b      | 2.03 ± 0.12 | 1.92 ± 0.08 | 1.00 ± 0.02 | 3.47 ± 0.12 | -   | G0V       | 9590                | 0.2–4     | 0.95 ± 0.03 | -0.3          | 1.11         |
| CoRoT-2b      | 3.31 ± 0.16 | 1.465 ± 0.029 | 1.74 ± 0.0281 | 9.102 ± 0.0281 | -   | K0V       | 5605                | 0.2–4     | 0.97 ± 0.06 | 0.902         | 1.25         |
| CoRoT-3b      | 21.66 ± 1.0 | 1.01 ± 0.07 | 4.25 ± 0.0577 | 9.277 ± 0.0577 | -   | F3V       | 6740                | 1.6–2.8   | 1.37 ± 0.2 | -0.02 ± 0.2 | 1.56         |
| CoRoT-4b      | 0.72 ± 0.08 | 1.19 ± 0.06 | 9.20 ± 0.09 | 8.456 ± 0.09 | -   | F0V       | 6890                | 0.2–4     | 1.10 ± 0.1 | 1.13 ± 0.1 | 1.15         |
| HAT-P-2b      | 8.04 ± 0.40 | 0.98 ± 0.40 | 5.63 ± 0.0677 | 8.962 ± 0.0677 | 0.5163 | F8        | 6620                | 2.7 ± 1.4 | 1.31 ± 0.1 | 0.12 ± 0.4 | 1.48         |

* Barge et al. (2008), * Alonso et al. (2008), * Deleuil et al. (2008), * Moutou et al. (2008), * Winn et al. (2007).

For all the CoRoT objects, $e = 0$ is assumed in the analysis of the light-curve and of the radial velocity data.

![Fig. 1. CoRoT-4b: evolution of the radius as a function of age. Solid line: standard cooling sequence of an isolated 0.72 M$\odot$ gaseous sphere with solar composition. Long-dashed line: irradiated case. Dash-dotted line: irradiated case with a 10 $M_\oplus$ water core. Red box: observational 1σ error bar.](image)

![Table 1. Transiting Exoplanets characteristics.](image)

4. Abnormally inflated planets

A significant fraction of transiting planets shows abnormally large radii (Mandushev et al. 2007) compared to predictions of
standard irradiated models (see e.g. Guillot 2008; Chabrier et al. 2009; Baraffe et al. 2009, and references therein). So far, two objects in the CoRoT sample belong to this category: CoRoT-1b and -2b. Comparison between models and observations for these two planets are shown in Figs. 2 and 3, respectively. The effect of irradiation alone does not allow to reproduce the observed radii of these strongly inflated objects. These observations thus confirm the existence of a missing mechanism in the models, which slows down the cooling and the contraction of the planet.

Several possibilities have been suggested to explain this puzzling property. Tidal heating due to circularization of the orbit, as originally suggested (Bodenheimer et al. 2001) and then rejected on the basis of a too short characteristic timescale compared with the age of the systems, might provide or at least participate to the lacking mechanism in some cases, provided tidal effects in the planet and in the star are properly taken into account (Jackson et al. 2008; Levrard et al. 2009; Ibgi & Burrows 2009). In a recent paper, however, Levrard et al. (2009) show that most transiting systems are not in a state of tidal equilibrium and thus that tidal circularization timescale estimates based on equilibrium tides are not correct. A major consequence of Levrard et al. (2009) calculations is that assuming zero-eccentricity, when not directly determined, in transiting light curve analysis, is not necessarily correct and such analysis should be redone carefully. Tidal heating might thus in some cases provide an extra source of energy during a significant fraction of the planet’s evolution. In the present models, we account for this source of energy, according to Hut (1981). Note that, since zero eccentricity is assumed for the analysis of all the CoRoT transit light-curves, we assume a small but finite value for $e$ in order to have a quantitative estimate of the effect of tidal heating on the bloated CoRoT planets. In the case of CoRoT-1b, an eccentricity $e = 0.02$ provides enough tidal heating to reproduce the observed radius within the error bars (dash-dotted curve in Fig. 2). For CoRoT-2b, an eccentricity of only a few percents does not provide enough tidal heating and larger values are required to reproduce the radius. We find that a value $e = 0.15$ yields a radius in agreement with observations.

Showman & Guillot (2002) suggested another heating mechanism in the deep interior of planets, originating from the strong winds generated at the planet’s surface. Their numerical simulations of atmospheric circulation produce a downward kinetic energy flux of about 1% of the absorbed stellar incident flux, which dissipates in the interior and slows down the planet’s contraction (Guillot & Showman 2002). Although the validity of this scenario is still debated, with various simulations producing different results (see e.g. Showman et al. 2008), it is worth exploring this issue. To test this effect, we have included an extra source of energy corresponding to 1% of the stellar impinging flux contribution in our evolutionary models, as done in Chabrier et al. (2004). In the case of CoRoT-1b (dashed curve in Fig. 2), the resulting theoretical radius is too large, by \~8%, compared to the observed value, while for CoRoT-2b, the same relative contribution yields too small of a radius. Note that for CoRoT-4b, analysed in Sect. 3, models including 1% of the stellar impinging flux as an extra source of energy also overestimate by 4% the radius compared to the observed value, except if the heavy element content of the planet is significantly larger than the inferred \~5% mass fraction. Enhanced atmospheric opacity in transiting planets (Burrows et al. 2007), although not excluded, remains so far too much of an ad-hoc suggestion to be examined in detail (see Baraffe et al. 2009, for a discussion). At any rate, the presently detected most inflated transiting planets, like Tres-4b and WASP-12, can not be explained even with such enhanced-opacity models (Guillot 2008; Baraffe et al. 2009).

An other mechanism, based on (inefficient) layered or oscillatory convection in some planet interiors, has been suggested by Chabrier & Baraffe (2007) and has been shown to possibly explain the abnormally large radii. If future follow-up observations of CoRoT-1b and CoRoT-2b confirm eccentricity values $e < 0.01$, such a mechanism will have to be considered with serious attention.

5. Massive substellar objects

The discovery of “super” Jupiters, with masses $\geq 10 M_J$ in close orbit to a central star, raises questions about their nature: planet or brown dwarf? CoRoT-3b (see Table 1) is a perfect example of such an ambiguity. Studies of low mass stars and brown dwarfs in young clusters suggest a continuous mass function down to $\sim 6 M_J$ (Caballero et al. 2007), indicating that the same formation process responsible for star formation can produce objects down to a few Jupiter masses. Indeed, analytical theories of star formation (Padoan & Nordlund 2004; Hennebelle & Chabrier 2008) show that gravoturbulent fragmentation of molecular clouds produces, with the same underlying processes, stars and brown dwarfs down to a few Jupiter-masses in numbers comparable to the observationally determined distribution. Brown dwarfs and planets thus overlap in mass, stressing the need for identification criteria enabling the distinction between these two types of astrophysical bodies. The presence of strongly non-solar atmospheric abundances, as observed in the atmosphere of the giant planets

**Fig. 2.** CoRoT-1b. Solid line: isolated solar composition giant planet. Long dashed line: irradiated case. Dash-dotted line: effect of tidal heating computed with $e = 0.02$. Short dashed line: solar composition case with 1% of the stellar flux transported deep into the planet interior.

**Fig. 3.** CoRoT-2b. Same legend as in Fig. 2.
of our Solar System, may provide signatures of a planetary formation process in a proto-planetary disk. Such a signature, however, is difficult both to observe and to characterize at the present time (Chabrier et al. 2007) and may not apply to irradiated planets, with radiatively stable outer layers. A more robust signature of the planet formation process, as expected from the core accretion model, is the presence of a significant amount of heavy material in the interior. Observed radii significantly smaller than predicted for solar or nearly-solar metallicity objects reveal the presence of such a significant average amount of heavy material; a major argument in favor of the core-accretion planet formation process. On the opposite, if a physical mechanism is missing in current planet cooling models, as discussed in Sect. 3, observed radii larger than predicted do not necessarily imply an absence or a small amount of heavy material. For such cases, the nature of the object remains ambiguous, if only based on the knowledge if its mean density.

In this section, we focus on the case of CoRoT-3b and examine whether its radius determination enables us to identify its very nature. As shown in Fig. 4, the observed radius of CoRoT-3b can be matched by the model of an irradiated brown dwarf of 21.6 $M_J$ with solar composition (long-dashed line). This is by itself an encouraging confirmation of the theoretical prediction of the age-mass-radius relationship in the brown dwarf regime (Chabrier & Baraffe 1997; Baraffe et al. 1998). Note that, given the small orbiting distance, the effects of irradiation are not negligible, even for such massive objects. Accounting for irradiation on the atmospheric profile, and thus on the object’s cooling history, is thus mandatory to provide consistent comparison between models and observations, when the radius is determined at this level of accuracy ($\sim 7\%$). The present radius error bars, however, are still too large to infer or exclude the presence of a significant amount of heavy material in the interior of this object. As done in Sect. 3, a maximum amount of heavy material can be determined, for the minimum theoretical radius allowed by the error box. We find an upper mass limit for the core of about 800 $M_\oplus$, i.e. a global maximum mass fraction $Z \lesssim 12\%$. We stress that this corresponds to the maximum enrichment compatible with the actual error bars. The possibility to have such a large amount of heavy material must be examined in the context of our current understanding of planet formation, within the framework of the core-accretion model. Following our previous analysis (Baraffe et al. 2008 and references therein), we can estimate the maximum amount of heavy material available in the proto-planetary disk for planet formation. According to current models of planet formation which include migration (Alibert et al. 2005), up to $\eta \sim 30\%$ of heavy elements contained in the protoplanetary disk can be incorporated into forming giant planets (Mordasini et al. 2008, 2009). The maximum mass of available heavy material that can be accreted to form planets is thus:

$$M_Z = \eta \cdot Z \cdot (f \cdot M_\star)$$

where $f \cdot M_\star$ is the maximum mass for a stable disk ($\lesssim 0.1 \, M_\star$) and $Z$ is the metal mass fraction of the star. For CoRoT-3b, which is orbiting a 1.37 $M_\odot$ F star with near solar metallicity, at most $M_Z \approx 270 \, M_\oplus$ of heavy material can thus be accreted to form the planet. This (admittedly crude) upper limit derived from current planet core accretion formation models yields a planet contraction consistent with today’s observations, as seen in Fig. 4. Note that, as discussed in Baraffe et al. (2008) for HAT-P-2b, this heavy material does not need to be accreted into one single object, as very massive planets, in particular short-period ones, may result from smaller planet collisions. Therefore, given the present uncertainties in the radius determination, neither the brown dwarf nor the planet possibility can be assessed or excluded for CoRoT-3b, whose nature remains ambiguous. A comparison between the predicted radius of a solar-metallicity brown dwarf (dashed line) and of a planet with the aforedetermined massive core, which represents only a $\sim 4\%$ metal enrichment (dash-dotted line) in Fig. 4, shows that a radius accuracy $\lesssim 3\%$ is required to resolve the ambiguity, according to the present models. In any event, this demonstrates the promising powerful diagnostic provided by mass-radius determinations to distinguish massive planets from low-mass brown dwarfs, providing adequate observational accuracy.

Among the few known massive planetary-mass objects, there is one example for which such a radius measurement provides the identification of its nature. This is the case of Hat-P-2b, a $8 \, M_\oplus$ mass object, closely orbiting an F type star (see Winn et al. 2007, and Table 1). As illustrated in Fig. 5, an irradiated brown dwarf model (long-dashed line) overestimates the radius by $\sim 5\%$. Models including a 340 $M_\oplus$ core mass\(^1\) can explain the measured radius (dash-dotted line). Although this amount of heavy material is about the limit of what is available for planet formation, according to current core-accretion models (as estimated from Eq. (1) for the HAT-P-2 system; see also Mordasini et al. 2009), the presence of such a metal enrichment ($Z \leq 15\%$) provides the simplest plausible explanation for the observed radius of HAT-P-2b, according to the present theory. As mentioned earlier, this 340 $M_\oplus$ core for HAT-P-2b should be seen as a rough estimate of the upper limit for the available heavy material in the system, but this analysis shows that the currently observed low radius of this object cannot be explained without a substantial enrichment. It would certainly be interesting to see whether planet models from other groups yield or not similar determinations. Given the fact that these various planet models share many common physics inputs (in particular the H/He and heavy element EOS), it would be surprising that they reach severely different conclusions. While keeping in mind the remaining uncertainties in planet cooling theory, the present analysis provides – with the parameters observed so far – a confirmation of the validity of the core-accretion model, and makes Hat-P-2b the first

\(^1\) Note that – as discussed in Sect. 2 – this amount of heavy material does not necessarily need to be in a core but could be distributed all over the planet.
confirmed \(8 M_{\oplus}\) genuine planet formed by core-accretion in a proto-planetary disk.

6. Conclusion

This work focusses on the modelling of the first confirmed transiting planets discovered by the CoRoT mission. We have distinguished three sorts of objects. First, planets whose radius can be explained by standard structure and evolution planetary models including the effect of irradiation from the parent star. For these objects, such as CoRoT-4b, there is no need to invoke extra physical mechanisms and, in that case, an upper limit can be inferred on their global heavy material content. The second category of objects is characterised by an abnormally large radius, with two examples in the current CoRoT sample. Ground-based surveys already found a significant fraction of such planets and CoRoT confirms this trend. We show that a small but finite eccentricity of \(\sim 0.02\) provides enough tidal energy dissipation to explain the radius of CoRoT-1b. For CoRoT-2b, a significantly larger value, \(e \sim 0.15\), is required. We emphasize that the fact that a zero eccentricity value is assumed in the light curve and radial velocity data analysis of current CoRoT planets, and of many other transiting objects. This hypothesis is based on the idea that (i) tidal circularisation of the orbit is the asymptotic equilibrium state of these planets; (ii) this circularisation occurs on short timescales compared to the system’s age. Two assumptions which have recently been shown not to be necessarily correct (Levrard et al. 2009). Eccentricity determinations of transiting planets, when unknown, should thus be redone in light of these results. Our estimates for the required eccentricity to lead to enough tidal dissipation to explain the observed radii could thus be compared with future follow-up observations/determinations.

The third category of planets includes the “massive” object CoRoT-3b, in the overlapping mass regime between brown dwarfs and giant planets. We show that the remaining large uncertainties in the radius determination (\(\sim 7\%\)) for this object do not allow a clear identification of its nature. For this object’s mass, present models predict less than 3% difference between the radius of a brown dwarf (solar composition) and of a planet with a realistic heavy material enrichment. By itself, the agreement between the observed radius and the theoretical predictions is a beautiful confirmation of the validity of these latter in the brown dwarf mass range, and brings confidence in the validity of the physics included in these models. Although our analysis is inconclusive for CoRoT-3b, given the present radius uncertainties, we show that the radius determination is discriminant in the case of Hat-P-2b, which is thus the first confirmation of the possibility to form massive planets by core accretion (possibly with subsequent collisions) up to \(m \gtrsim 8 M_{\oplus}\). Finally, our analysis shows that, according to the present models, a typical \(\lesssim 5\%\) accuracy on the radius determination must be achieved in future space-based or ground based transit detections to clearly distinguish planets from brown dwarfs in their overlapping mass domain. We stress, of course, that the conclusions of the present paper are based on presently published observational error bars and on our present theoretical description of planet structure and evolution, still hampered by many uncertainties. The present results thus need to be confronted to e.g. future COROT or KEPLER detections in order to assess their reliability.

References

Alibert, Y., Mordasini, C., Benz, W., & Winisdoerffer, C. 2005, A&A, 434, 343
Alonso, R., Auvergne, M., Baglin, A., et al. 2008, A&A, 482, L21
Baraffe, I., Chabrier, G., Allard, F., & Hauschildt, P. H. 1999, A&A, 337, 403
Baraffe, I., Chabrier, G., Barman, T. S., Allard, F., & Hauschildt, P. H. 2003, A&A, 402, 701
Baraffe, I., Chabrier, G., & Barman, T. 2008, A&A, 482, 315
Baraffe, I., Chabrier, G., & Barman, T. 2009. Reports on Progress in Physics, Submitted
Barge, P., Baglin, A., Auvergne, M., et al. 2008, A&A, 482, L17
Barman, T. S., Hauschildt, P. H., & Allard, F. 2001, ApJ, 556, 885
Bodenheimer, P., Lin, D. N. C., & Mardling, R. A. 2001, ApJ, 548, 466
Boley, A. C., Durisen, R. H., Nordlund, A., & Lord, J. 2007, ApJ, 665, 1254
Boss, A. P. 1997, Science, 276, 1836
Burrows, A., Sudarsky, D., & Hubbard, W. B. 2003, ApJ, 594, 545
Burrows, A., Hubeny, I., Budaj, J., & Hubbard, W. B. 2007, ApJ, 661, 502
Caballero, J. A., Bejar, V. J. S., Rebolo, R., et al. 2007, A&A, 470, 903
Chabrier, G., & Baraffe, I. 1997, A&A, 327, 1039
Chabrier, G., & Baraffe, I. 2007, ApJ, 661, L81
Chabrier, G., Barman, T., Baraffe, I., Allard, F., & Hauschildt, P. H. 2004, ApJ, 603, L53
Chabrier, G., Baraffe, I., Selosse, F., et al. 2007, in Protostars and Planets V, ed. B. Reipurth, D. Jewitt, & K. Keil, 623
Chabrier, G., Baraffe, I., Leconte, J., Gallardo, J., & Barman, T. 2009, ed. E. Stempels, AIP Conf. Ser., 1094, 102
Deleuil, M., Deeg, H. J., Alonso, R., et al. 2008, A&A, 491, 889
Dullemond, C., Durissen, R., & Papaloizou, J. 2009, in Structure Formation in Astrophysics, ed. G. Chabrier
Fortney, J. J., Marley, M. S., & Barnes, J. W. 2007, ApJ, 668, 1267
Guillot, T. 2008, Phys. Scr. T. 130, 014023
Guillot, T., & Showman, A. P. 2002, A&A, 385, 156
Guillot, T., Burrows, A., Hubbard, W. B., Lumme, J. I., & Saumon, D. 1996, ApJ, 459, L35
Hennebelle, P., & Chabrier, G. 2008, ApJ, 684, 395
Hut, P. 1981, A&A, 99, 126
Ibgui, L., & Burrows, A. 2009, ApJ, 700, 1921
Levrard, B., Winisdoerffer, C., & Chabrier, G. 2009, ApJ, 692, L9
Mandushev, G., O’Donovan, F. T., Charbonneau, D., et al. 2007, ApJ, 667, L195
Mordasini, C., Alibert, Y., Benz, W., & Naef, D. 2008, ed. D. Fischer, F. A. Rasio, S. E. Thorsett, & A. Wolszczan, ASP Conf. Ser., 398, 235
Mordasini, C., Alibert, Y., & Benz, W. 2009, A&A, 501, 1139
Moutou, C., Bruntt, H., Guillot, T., et al. 2008, A&A, 488, L47
Padovan, P., & Nordlund, Å. 2004, ApJ, 617, 559
Rafikov, R. R. 2005, ApJ, 621, L69
Rafikov, R. R. 2007, ApJ, 662, 642
Saumon, D., Chabrier, G., & van Horn, H. M. 1995, ApJS, 99, 713
Showman, A. P., & Guillot, T. 2002, A&A, 385, 166
Showman, A. P., Menou, K., & Cho, J. Y.-K. 2008, ed. D. Fischer, F. A. Rasio, S. E. Thorsett, & A. Wolszczan, ASP Conf. Ser., 398, 419
Thompson, S., & Launus, H. 1972, Technical Report Technical Report SC-RR-61 0714, Sandia National Laboratories
Winn, J. N., Johnson, J. A., & See, K. M. G., et al. 2007, ApJ, 665, L167

Fig. 5. Hat-P-2b. Long dashed line: cooling sequence of an irradiated \(8 M_{\oplus}\) brown dwarf. Dash-dotted line: irradiated case with a \(340 M_{\oplus}\) core. Red box: observational 1σ error bar.