Sealed operation, and circulation and purification of gas in the HARPO TPC

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Abstract. HARPO is a time projection chamber (TPC) demonstrator of a gamma-ray telescope and polarimeter in the MeV-GeV range, for a future space mission. We present the evolution of the TPC performance over a five month sealed-mode operation, by the analysis of cosmic-ray data, followed by the fast and complete recovery of the initial gas properties using a lightweight gas circulation and purification system.

1 Introduction

High-energy astronomy is hindered by the sensitivity gap between the energy ranges over which the Compton and the pair telescopes are efficient, making the γ-ray sky in the range 0.1 – 100 MeV almost unexplored. Furthermore no polarimeter with a decent sensitivity above 1 MeV has ever been flown in space.

On the low-energy side of the gap, a number of projects are aiming at an improved sensitivity with respect to the COMPTEL telescope [7]. Of these, those using stacks of silicon wafers without any tungsten converter can also detect pair conversions [1–3] with an angular resolution of ≈ 1.5° @ 100 MeV, a significant improvement over that of Fermi-LAT (≈ 6° @ 100 MeV [4]). These stacks can have some sensitivity to polarization if the wafers are thin enough and provided that matching can be performed in the first wafer after the conversion wafer, i.e. that each of the hits in the x direction can be assigned to one of the two hits in the orthogonal (y) direction. In subsequent wafers, the azimuthal information carried by the pair is washed out by multiple scattering [5].

In the quest for even thinner wafers, it is tempting to use an even less dense detector such as a gas detector. With a spatial resolution of 100 µm, and with an optimal tracking in the presence of multiple scattering provided by a Kalman filter, an angular resolution of ≈ 0.4° @ 100 MeV can be achieved [6]. This provides a point-like source sensitivity as low as 10−4 MeV/(cm²s) @ 100 MeV with a 10 kg gas time projection chamber (TPC), while that of COMPTEL plateaued at ≈ 10−2 MeV/(cm²s) [7]. In such gas detectors the dilution of the azimuthal information carried by the pair is close to unity, so that the sensitivity to polarization is excellent [8]. Finally track matching can easily be performed by comparing the distribution of the energy deposited along the tracks in the x and y directions thanks to the large fluctuation in the local deposition of energy along the track [9]. These unique features come at the cost of the need to maintain a good-quality gas in which the collection of the ionizing electrons is not affected by impurities with a large electron attachment. For example the EGRET gas (99.5% neon, 0.25% argon, and 0.25% ethane) became contaminated by the breakdown of ethane by spark chamber firings and to a small degree by residual outgassing. The five planned gas exchanges were carried out once per year, after which the lack of fresh gas became a serious issue [10]. These concerns have triggered the development of detectors built out of low outgassing rate materials. For example, the design of the GEMS X-ray polarimeter TPC involves gas electron multipliers (GEM) made of copper clad liquid crystal polymer (LCP), and a dielectric structure of polyetheretherketone (PEEK) on a gold-plated titanium frame. This gives it an expected lifetime of 23 years [11].

We report here on a gas purity study performed with the HARPO detector, a TPC prototype of a γ → e⁺e⁻ telescope built with techniques much better suited to our actual budget and that we have studied in the laboratory with radioactive sources [12] and with cosmic rays [13], and that we have tested recently [14] in a γ-ray beam.

2 Description of the HARPO TPC

HARPO is a 30 cm cubic TPC surrounded by 6 scintillator plates, each equipped with two wavelength shifter bars and two photomultipliers (PMTs) for trigger and background rejection. The gas container is an aluminum cylinder designed to be operated from few mbar up to 5 bar in sealed mode (Fig. 1). With a mixture of 95% of argon 5.0 and 5% of iC₄H₁₀ 3.5 and at a 220 V/cm drift electric field, the electron drift velocity is about 3 cm/µs. The TPC endplate is a "hybrid" MicroPattern Gas Detector (MPGD)
composed of two GEM foils [15] with 2 mm spacing above a 128 $\mu$m amplification gap bulk-micromegas [16]. The anode collection plane is segmented in $2 \times 288$ strips in the two orthogonal $x, y$ directions and the signal is read by an electronic chain composed of two T2K/TPC Front-End cards [17], two FEMINOS back-end cards and a Trigger Clock Module [18]. A simple coincidence between scintillators is used here for triggering on cosmic rays.

3 Description of the circulation and purification system (CAPS)

We used a circulator composed of a simple turbine (Fig. 2, left). Full schematics of the circulation system is shown in Fig. 3 The gas was lead through an Oxisorb purification cartridge that can absorb up to 0.5 L H$_2$O and 0.1 L O$_2$, with output purity as low as < 5 ppb for O$_2$ and 30 ppb for H$_2$O [19]. We measured the gas flow with a F-111B Bronkhorst mass flowmeter (Fig. 3, I10) that was calibrated for our gas mix at 2 bar and 20 °C. We monitored the gas pressure and temperature with a Keller EV-120/PA(A)-33X sensor (Fig. 3, I4 and I5) with ranges of 0 − 50 °C and 0 − 10 bar and resolutions of 0.01 °C and 1 mbar. The outgassing areas inside the vessel amount to 0.82 m$^2$ Aluminum, 3.50 m$^2$ PVC, 0.92 m$^2$ PCB, 1.63 m$^2$ PVT, for a gas volume $V_{\text{det}} = 84$ L.

We tested the gas circulation system with a helium leak detector (ASM 310) down to $3.3 \times 10^{-8}$ mbar·L/s, after which we opened the connection between the CAPS and the detector. This caused a pressure drop from 2015 mbar to 1960 mbar at 27 °C, with a flow of 1.5 L/h. On July 20, 2015, we opened the Oxisorb purifier in vacuum which lead to a new pressure drop to 1932 mbar when the connection to the detector was opened. The CAPS then ran continuously, with a flow through the filter of $f_{\text{circ}} = 1.4$ L/h. In these conditions we expect the water and oxygen content in the gas, $C$, to evolve following a law of the form:

$$C = C_0 e^{-t/\tau_{\text{purif}}} + K_{\text{cont}} f_{\text{circ}},$$

with $\tau_{\text{purif}} = V_{\text{det}}/\epsilon_{\text{purif}} f_{\text{circ}}$, where $C_0$ is the original content and $K_{\text{cont}}$ is the contamination rate from leaks and outgassing. In case the purification efficiency of the filter is $\epsilon_{\text{purif}} = 1$, the purification rate is equal to $\tau_{\text{purif}} = 60$ h.

The Oxisorb purifier color ranges from blue to dark brown with the quantity of H$_2$O and O$_2$ absorbed. At the beginning of the purification, we observed a purple ring of 3 mm thickness growing 3 mm per day during the first 4 days. It stabilised at 12 mm after a few days (Fig. 2). This indicates that the cartridge was never saturated.

Several gas samples were taken for analysis, which lead to density drops (lines 1, 2, and 3 on Fig. 7). Before and after every operation that affected the gas in the vessel, data was taken with the TPC for consistency checks.

4 Data taking and analysis

We took cosmic-ray data once per week from February to July, with the detector set so that the drift field was vertical. Most runs were one hour long, recording a few 10,000 events. After starting the purification on July 20
we took data more frequently. The high voltage was only turned on when taking data. Due to the vertical detector orientation, most of the recorded tracks are approximately aligned with the drift direction. Figure 4 shows the X, t and Y, t projections for a typical event. We select events with a single track in each projection, with a high momentum muon selection based on the $\chi^2$ cut of a linear fit. For each time bin $t_i$ (30 ns $\approx$ 1 mm), we obtain the charge deposited by a track, $Q(t_i)$, summing over all the channels associated with the track. After an angle-effect correction using the direction $\vec{u}$ of the track,

$$\frac{dE}{dx}(t_i) = Q(t_i) \times \vec{u}$$  \hspace{1cm} (2)

we obtain the $dE/dx(t_i)$ in ADC units, that is up to an overall factor that includes the gain of the amplification of both the MPGDs and that of the readout electronics. The resulting distribution (Fig. 5) is also affected by threshold and saturation effects. To mitigate them, we estimate the most probable value (MPV, solid line in Fig. 5) of the charge for each time bin with a Landau-function fit. From this distribution we extract:

**The drift velocity.** As the tracks are selected to traverse the full TPC length, the drift-time duration provides a measurement of the drift velocity.

**The gain.** The height of the distribution at short drift times is proportional to the amplification gain.

**The attenuation.** The signal loss along the drift from electron capture is visible in Fig. 5. We fit the central part of the distribution, which is less affected by the threshold and saturation effects, with an exponential function.

These measurements are affected by systematic effects. Since the detector configuration was kept unchanged throughout the data taking, they cancel in the relative value compared to a reference run, chosen to be run 2008 (Feb. 18th), i.e. right after the fresh gas fill. Figure 6 shows the relative charge distribution as a function of the drift distance for several runs taken at a few weeks interval.

**Table 1.** Gas content from mass spectrometry (volume/volume). Note that the set-up was not sensitive to H2O.

| Sample # | Compound | 1 July 08 | 2 Aug. 27 | 3 Sept. 17 |
|----------|----------|-----------|-----------|------------|
|          | C4H10    | %         | ppm       | ppm        |
|          | O2       | 180       | <20       | <20        |
|          | CO       | 190       | 250       | 130        |
|          | CO2      | 120       | 160       | 130        |
|          | N2       | 620       | 890       | 850        |

**Figure 6.** Charge as a function of the drift distance, for three runs, 6 weeks apart, normalised to the reference run 2008.

**Figure 4.** A cosmic-ray event in HARPO. The “maps” show the charge distributions in the projection X-time and Y-time.

**Figure 5.** Normalised charge as a function of the drift time for a 6000 s cosmic-ray run.
inside the almost closed field cage. We also observe a decrease of the drift velocity, which cannot be associated with the oxygen contamination and which is recovered after purification, and therefore that can be attributed to water contamination.

After purification, the drift velocity was measured to be $6.0 \pm 0.5\%$ higher than its initial value. From the Garfield simulation, this can be explained by the cumulative effects of the reduced gas pressure from 2.1 to 1.9 bar (+4%) and of the measured decrease of the isobutane content from 5.1% to 4.5% (+3%).

7 Conclusion

We have successfully operated a sealed TPC over five months, with a clear but manageable gas degradation. We used the data from cosmic rays in the TPC to monitor the relevant parameters of the gas (drift velocity, gain, attenuation). We then purified the gas with a simple, low-power circulation system and observed a fast and complete recovery of the initial parameters. If operated routinely, for example on a space mission, the system should therefore enable a stable performance for several years.

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References

[1] “ASTROGAM”, proposal submitted for the ESA M4 Mission Programme, Jan. 2015, http://astrogam.iaps.inaf.it/
[2] A. A. Moiseev et al., arXiv:1508.07349 [astro-ph.IM].
[3] X. Wu et al., Proc. SPIE 9144, 91440F (2014) [arXiv:1407.0710 [astro-ph.IM]].
[4] M. Ackermann et al. [Fermi-LAT Collaboration], Astrophys. J. Suppl. 203 (2012) 4 [arXiv:1206.1896 [astro-ph.IM]].
[5] D. Bernard, The gamma-ray sky with astrogam, Second ASTROGAM Workshop, March 2015, Paris.
[6] D. Bernard, Nucl. Instrum. Meth. A 701, 225 (2013) [Erratum-ibid. A 713, 76 (2013)] [arXiv:1211.1354 [astro-ph.IM]].
[7] V. Schönfelder, New Astr. Rev. 48 (2004) 193.
[8] D. Bernard, Nucl. Instrum. Meth. A 729 (2013) 765, [arXiv:1307.3892 [astro-ph.IM]].
[9] D. Bernard, 12th Pisa Meeting on Advanced Detectors, Isola d’Elba, Italy, May 2012, Nucl. Instrum. Meth. A 718, 395 (2013), [arXiv:1210.4399 [astro-ph.IM]].
[10] J. A. Esposito et al., Astrophys. J. Suppl. 123 (1999) 203.
[11] J. E. Hill et al., Proc. SPIE 8859, San Diego, United States, Aug. 2013, 88590O.
[12] P. Gros et al., TIPP14, 2-6 June 2014 Amsterdam, proceedings PoSTIPP2014/p133.
[13] D. Bernard et al., SPIE 9144, June 2014, Montréal, Canada, 91441M; arXiv:1406.4830 [astro-ph.IM].
[14] S. Wang et al., TPC2014, Dec. 2014 Paris, J.Phys.Conf.Ser. 650 (2015) 012016, arXiv:1503.03772 [astro-ph.IM].
[15] F. Sauli, Nucl. Instrum. Meth. A 386, (1997) 531.
[16] I. Giomataris et al., Nucl. Instrum. Meth. A 560, (2006) 405.
[17] P. Baron et al., IEEE Trans. Nucl. Sci. 55 (2008) 1744.
[18] D. Calvet, IEEE Trans. Nucl. Sci. 61 (2014) 675.
[19] Oxisorb, air liquide, ref 123342.
[20] R. Veenhof et al., http://garfieldpp.web.cern.ch/garfieldpp/