Tidal Tails around the Outer Halo Globular Clusters Eridanus and Palomar 15

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Abstract

We report the discovery of tidal tails around the two outer halo globular clusters, Eridanus and Palomar 15, based on g-band images obtained with DECam at the CTIO 4 m Blanco Telescope. The tidal tails are among the most remote stellar streams currently known in the Milky Way halo. Cluster members have been determined from the color–magnitude diagrams and used to establish the radial density profiles, which show, in both cases, a strong departure in the outer regions from the best-fit King profile. Spatial density maps reveal tidal tails stretching out on opposite sides of both clusters, extending over a length of ~760 pc for Eridanus and ~1160 pc for Palomar 15. The great circle projected from the Palomar 15 tidal tails encompasses the Galactic Center, while that for Eridanus passes close to four dwarf satellite galaxies, one of which (Sculptor) is at a comparable distance to that of Eridanus.

Key words: globular clusters: general – globular clusters: individual (Eridanus, Palomar 15)

1. Introduction

It is widely accepted that large galaxies such as the Milky Way formed through accretion and merger of numerous protogalactic fragments (e.g., Searle & Zinn 1978; Blumenthal et al. 1984). A significant fraction of the Milky Way’s globular cluster population is thought to have been acquired by this process, and it is believed that a large portion of the current halo stellar mass may have been donated by their host dwarf galaxies via tidal disruption and mass loss (Mackey & Gilmore 2004; Forbes & Bridges 2010).

There are two types of stellar structures that we expect to see around Galactic globular clusters. One can occur around accreted clusters and represents the remnant of the disrupted parent dwarf galaxy (see Olszewski et al. 2009 and Kazma et al. 2016). In extreme cases, we might see a cluster embedded in an extended stellar stream. This is the case in the halo of M31 (e.g., Mackey et al. 2010) and in our own Galaxy where a number of globular clusters are associated with the disrupting Sagittarius dwarf galaxy (e.g., Law & Majewski 2010).

The second type consists of narrow stellar streams arising from the dynamical evolution of the cluster itself in the external tidal field of the Milky Way. Many studies have investigated the presence of tidal tails associated with globular clusters (e.g., Grillmair et al. 1995; Leon et al. 2000). It has been found that some globular clusters have stellar distributions that significantly differ from a King profile (King 1962), extending beyond the nominal tidal radius (Grillmair et al. 1995; Carullo-Bello et al. 2012).

Recently, it has been discovered that the tidal tails of some globular clusters, such as Palomar 5 (Odenkirchen et al. 2001; Grillmair & Dionatos 2006) and NGC 5466 (Belokurov et al. 2006a; Grillmair & Johnson 2006), extend over several hundred parsecs in physical length. In this context, it is also interesting to note that many narrow streams, like the Orphan (Belokurov et al. 2006b; Grillmair 2006; Grillmair et al. 2015) and Phoenix streams (Balbinot et al. 2016), exist in the inner Milky Way halo. These likely originate from completely disrupted globular clusters (Newberg et al. 2010; Martin et al. 2014).

The main point is that the Galactic tidal field, disk and bulge shocks, and two-body relaxation can all affect the outer structures of globular clusters in ways that we do not fully understand. Adding new examples of globular clusters with tidal tails, particularly at large Galactocentric distances, thus gives us additional information about the frequency of this phenomenon and can help to probe the outermost parts of the Galaxy. The properties of such streams can further help to constrain the dark matter distribution in the Galactic halo. For example, the gaps in the tidal tails of Palomar 5 may tell us about the dark matter sub-halos predicted in ΛCDM cosmology (Ngan & Carlberg 2014).

So far, most studies searching for tidal structures around Milky Way globular clusters have been restricted to relatively nearby targets (e.g., Grillmair & Carlin 2016) with the one exception of Palomar 14 (Sollima et al. 2011). In this Letter, we report the discovery of stellar substructures around two distant halo globular clusters, Eridanus (RGC = 95.0 kpc) and Palomar 15 (Pal 15; RGC = 38.4 kpc).

2. Data Analysis

We used the Dark Energy Camera (DECam) at the CTIO 4 m Blanco Telescope to obtain deep imaging around Eridanus and Pal 15. This instrument comprises a 62 CCD mosaic that spans a 3 deg² field of view and has a pixel scale of 0.′′27 (Flaugher et al. 2015). For Eridanus, we obtained 5 dithered 900 s g-band images and 11 dithered 600 s i-band images on 2014 February 25–27 (average seeing ~1′′11). For Pal 15, we obtained five dithered 360 s g-band images and five dithered 360 s i-band images on 2013 September 24–26 (average seeing ~1′′12). The nights were part of the observing programs 2014A-0621 for Eridanus and 2013B-0617 for Pal 15—PI: D. Mackey. The raw images were preprocessed with the DECam Community Pipeline (Valdes et al. 2014), including application of the astrometric solution. We used the resampled images (individual frames) and their corresponding weight maps for our study.

SExtractor (Bertin & Arnouts 1996) and PSFEx (Bertin 2011) were employed for source detection and PSF photometry. Star/galaxy separation was performed using the method described in Koposov et al. (2015), which adopts
For a given target we constructed the \((g - i)_0, g_0\) color-magnitude diagram (CMD) using all stellar objects within \(r < 3\)' from the nominal cluster center and fit a fiducial line representing the locus of cluster members in the CMD. The fiducial line for each cluster was determined empirically. The CMDs for the inner 3', which are dominated by cluster members, were subdivided into intervals of \(g_0\) and the 3\(\sigma\)-clipped median \((g - i)_0\) colors were determined. A continuous curve through the pairs of mean \(g_0\) and the corresponding median color for each interval then define the fiducial line for the cluster members. This fiducial line was then used to calculate a weight \(w\) for each star detected across the field. The weight quantifies the likelihood of a star being a cluster member based on its distance from the fiducial line along the color axis. The weight value was calculated using the Gaussian distribution \(N(x, \mu, \sigma)\) centered on the fiducial line (normalized to have \(w = 1.0\) at the center), with \(\sigma\) corresponding to the uncertainty in \((g - i)_0\) at the \(g_0\) value, which were taken from the photometric errors generated by SExtractor.

Stars within a 2\(\sigma\) range from the fiducial line in color (that is, 
\[ w = N((g - i)_0, (g - i)_{0,\text{fit}}), \sigma(g - i)_0 > 0.135 \]) were considered to have a high probability of being related to the target cluster. We will henceforth call this selection “member stars” although it still contains some contamination from foreground/background objects (henceforth “background”) that happen to lie near the cluster population in CMD. To minimize this contamination while maximizing the signal due to cluster member stars, we set a threshold for stars. The instrumental magnitudes were calibrated by cross-matching with the Pan-STARRS1 StackObjectThin catalog (Chambers et al. 2016) and then de-reddened using the Schlegel et al. (1998) dust maps with the extinction coefficients from Schlaufly & Finkbeiner (2011). The inferred distances are consistent with those given in Harris (1996).
members, we further restricted our selection to an interval in $g_0$ defined to cover the region on the CMD with the greatest contrast between “cluster” and “background” stars. In addition, we ensured that the faint end of this interval was 0.5 mag above the 50% photometric completeness limit.

Radial density profiles and spatial density maps were generated to investigate the distribution of the member stars (Figures 1 and 2). For the radial profile (upper left panel), we binned our field of view in concentric annuli with logarithmic spacing. Poisson statistics were assumed for calculating the uncertainty in each bin. The best-fitting King profile (King 1962) was determined by using a python library LIMFIT (Newville et al. 2014) after subtracting the background level. The background level was fixed from the total (cluster + background) density profile by determining the median annular density after 3σ-clipping.

Two-dimensional density maps (bottom panels) were constructed by binning the DECam field into cells of size $0.33 \times 0.33$ for Eridanus and $0.59 \times 0.59$ for Pal 15, and then smoothing with a Gaussian kernel of width $\sigma = 0.8$ for Eridanus and $\sigma = 1.47$ for Pal 15. The bin size and sigma were chosen after testing various combinations to provide the maximum contrast for the tidal features. The median stellar density of the field (=background), beyond 3 tidal radii from the cluster, is $5 \text{ star arcmin}^{-2}$ ($\sigma = 2.9 \text{ star arcmin}^{-2}$) for Eridanus and $7 \text{ star arcmin}^{-2}$ ($\sigma = 3.1 \text{ star arcmin}^{-2}$) for Pal 15 before smoothing. We generated a background map for each cluster in the same manner by using the corresponding subsample of background objects. After this map was normalized to the median 3σ-clipped density of the $w > 0.135$ map determined above, it was subtracted to remove any large-scale gradients or instrumental artifacts across the field. Local rms values were measured for each of the 62 CCDs. The median was considered to be the background rms and used to scale the background-subtracted spatial density map (bottom left panel).

3. Eridanus

Eridanus, at $R_{\text{GC}} = 95.0 \text{kpc}$ (Harris 1996), is one of the most distant Galactic globular clusters known. It has been classified as a “young” halo cluster, suggesting that it may have originated in an external satellite galaxy and been accreted into the Galactic halo.
The dynamical evolution of Eridanus in the tidal field of the Milky Way.

4. Palomar 15

Pal 15 is an outer halo globular cluster at $R_{GC} = 38.4$ kpc (Harris 1996). Although it is classified as an "old" halo cluster, Pal 15 is a good accretion candidate by virtue of its location at a large Galactocentric distance (Mackey & Gilmore 2004). Our best-fitting King profile has a core radius $r_c = 0.25 \pm 0.12$ (6.6 $\pm$ 3.1 pc) and a tidal radius $r_t = 3.17 \pm 0.76$ (83.0 $\pm$ 20.0 pc), with the concentration index $c = \log(r_t/r_c) = 1.10 \pm 0.23$. These values agree well with previous measurements ($r_c = 0.25$ and $r_t = 3.15$; Harris 1996). The radial profile exhibits an excess of stars outside $r \sim 1.82$ (log($r$) $\sim$ 0.26), which continues beyond the nominal tidal radius $r > r_t = 3.17$ and follows a power law with an index of $c = -1.64 \pm 0.16$ (azimuthally averaged).

The spatial density map and the corresponding contour map (lower panels in Figure 1) reveal two prominent tidal tails extending considerably beyond the tidal radius (dashed circle in the contour map). These structures show no correlation with the background map. Tail 1 extends $\sim$18$'$ ($\sim$480 pc) from the cluster center in the north–east direction at a position angle of PA = 40$^\circ$ $\pm$ 5$^\circ$. Tail 2 extends $\sim$11$'$ ($\sim$280 pc) in the south–west direction (PA = 220$^\circ$ $\pm$ 10$^\circ$). The fractional overdensity of cluster stars is 29.4% and 31.6%, respectively, in the distance interval 2.5 $r_c < r < 3 r_c$. The alignment of the features is close to axisymmetric, suggesting that they are likely following the orbit of Eridanus. Given the alignment and the narrow width compared to the size of the cluster, it would seem likely that the features are tidal tails resulting from the dynamical evolution of Eridanus in the tidal field of the Milky Way.

| ID      | $r_c$ (pc) | $r_t$ (pc) | $c$        | $l_{\text{size1}}$ (pc) | $l_{\text{size2}}$ (pc) | (angle$_{\text{mc}}$) |
|---------|------------|------------|------------|------------------------|------------------------|-------------------|
| Eridanus | 6.6 $\pm$ 3.1 | 83.0 $\pm$ 20.0 | 1.10 $\pm$ 0.23 | $\sim$480 | $\sim$280 | $\sim$180 |
| Pal 15  | 18.4 $\pm$ 2.0 | 68.1 $\pm$ 14.7 | 0.57 $\pm$ 0.11 | $\sim$780 | $\sim$380 | $\sim$170 |

Note. (1) indicates Tail 1, (2) indicates Tail 2, and (angle$_{\text{mc}}$) indicates the angle between two tails.

We have discovered tidal tails extending from the two remote Milky Way globular clusters Eridanus and Pal 15. The narrowness of the tails compared to the cluster size, along with their symmetric orientation on either side of the cluster centers indicates that they are due to the loss of cluster members to the Galactic tidal field, as seen for several other globular clusters such as Pal 5 and NGC 5466 and in numerical simulations (Combes et al. 1999; Capuzzo Dolcetta et al. 2005). Moreover, the power-law slope for each tail region (PA$\pm$10$^\circ$) shows $\gamma = -1.20 \pm 0.19$ and $\gamma = -1.25 \pm 0.16$ for Eridanus, $\gamma = -1.24 \pm 0.12$ and $\gamma = -1.26 \pm 0.12$ for Pal 15, which are similar values as in the case of Pal 5 (Odenkirchen et al. 2003) and steeper than constant density.

Table 1 contains the estimated basic parameters for the tidal tails. Both pairs exhibit significant extent beyond the tidal radius of their respective cluster—the tails of Eridanus span $\sim$760 pc in total, while those for Pal 15 trace $\sim$1160 pc.

The tidal tails of Pal 15 may well extend beyond the edge of the DECam field of view. In this case, our study provides a lower limit on their length, and some additional off-field data will be required to test the true extent of Pal 15’s tidal structure. In addition, the tails appear curved, especially if the fragment near the southern edge of our Pal 15 field is considered to be part of the structure. If confirmed, this could hold information about the underlying Galactic potential or it may be simply related to projection effects (Combes et al. 1999).

For Eridanus, the tidal tails are well confined to the inner part of our DECam field. This suggests that we have effectively found their full extent, unless the star densities in the tails are significantly reduced.
Canes Venatici I, Canes Venatici II, Fornax, and Sculptor provide a rough estimate for the cluster uncertainty estimated for each tail. Since the tidal tails are expected to be projected along the orbital path of the cluster (Combes et al. 1999; Capuzzo Dolcetta et al. 2005), this great circle can provide a rough estimate for the cluster’s orbit. It is notable that the potential orbit of Eridanus passes close to the dwarf galaxies Canes Venatici I, Canes Venatici II, Fornax, and Sculptor. Sculptor (R$_{GC} = 86$ kpc; McConnachie 2012) has a comparable distance to that of Eridanus. Since Eridanus is classified as a "young" halo cluster based on its metallicity and horizontal branch morphology (Mackey & van den Bergh 2005), and indeed exhibits a CMD consistent with an age up to ~2 Gyr younger than the oldest Milky Way globular clusters (Stetson et al. 1999), this possible association to Sculptor and/or another of the three satellites might provide additional evidence of Eridanus’s extragalactic origin (see also Keller et al. 2012). Even so, it is striking that despite several indications that both Eridanus and Pal 15 might be accreted, and our data being sensitive enough to detect low surface brightness features at their distances, no clear evidence for stellar debris from parent dwarf galaxies was observed. Carballo-Bello et al. (2015) suggested that Eridanus may be associated with the Monoceros ring based on the modeled orbit of this structure by Penarrubia et al. (2005); however, the great circle of Eridanus disagrees with this modeled orbit, suggesting no clear association between them.

The great circle defined by the tidal tails of Pal 15 passes close to Ursa Major II, which also has a comparable distance of R$_{GC} = 38$ kpc (McConnachie 2012), suggesting possible association. We further notice that the potential orbits of Eridanus and Pal 15 extend to relatively high Galactic latitudes, indicating that the clusters may be on plunging orbits relative to the Galactic disk. Assuming the orbits are sufficiently eccentric to pass through the inner region of the Milky Way, the clusters could experience periodic disk and/or bulge shocks. Since these processes are known to be destructive (Gnedin & Ostriker 1997; Gnedin et al. 1999), this could explain why both clusters exhibit tidal tails. A prime example is Pal 5, which suffers this type of tidal shock and is predicted to be destroyed at its next disk crossing (Dehnen et al. 2004).

In the case of Eridanus (R$_{GC} = 95.0$ kpc), we have discovered a new tidal stream in an extremely remote part of the Galaxy which is poorly understood and where no other narrow streams are known. Studying Eridanus and its tails in detail may lead us to a better understanding of the gravitational potential in the extreme outskirts of the Galaxy. A key aspect of this will be understanding whether the tails arise from the action of the tidal field at this large Galactocentric distance, or whether Eridanus is on a very eccentric orbit such that it passes through the inner Milky Way, and the tails are more likely due to the action of the tidal field at smaller radii.

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Figure 3. Great circle along the direction of the tidal tails of Eridanus (blue line) and Pal 15 (red line). Dashed lines in the same color indicate the corresponding uncertainty range. The two stars are Eridanus (blue) and Pal 15 (red). Milky Way globular clusters listed in Harris (1996) are marked as dots in blue for “young” halo clusters, red for “old” halo clusters, and gray for unclassified (Mackey & van den Bergh 2005). Major Milky Way dwarf satellite galaxies listed in McConnachie (2012) are marked with black triangles. Vertical and horizontal gray dashed lines indicate 0° of Galactic longitude and latitude, respectively.
