Quasars, their host galaxies, and their central black holes

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ABSTRACT

We present the final results from our deep HST imaging study of the host galaxies of radio-quiet quasars (RQQs), radio-loud quasars (RLQs) and radio galaxies (RGs). We describe and analyze new WFPC2 $R$-band observations for 14 objects which, when combined with the first tranche of HST imaging reported in McLure et al. (1999), provide a complete and consistent set of deep, red, line-free images for statistically-matched samples of 13 RQQs, 10 RLQs and 10 RGs in the redshift band $0 < z < 0.25$. We also report the results of new deep VLA imaging which has yielded a 5 GHz detection of all but one of the 33 AGN in our sample.

Careful modelling of our images, aided by a high dynamic-range point-spread function, has allowed us to determine accurately the morphology, luminosity, scale-length and axial ratio of every host galaxy in our sample. Armed with this information we have undertaken a detailed comparison of the properties of the hosts of these 3 types of powerful AGN, both internally, and with the galaxy population in general.

We find that spheroidal hosts become more prevalent with increasing nuclear luminosity such that, for nuclear luminosities $M_V < -23.5$, the hosts of both radio-loud and radio-quiet AGN are virtually all massive ellipticals. Moreover, we demonstrate that the basic properties of these hosts are indistinguishable from those of quiescent, evolved, low-redshift ellipticals of comparable mass. This result rules out the possibility that radio-loudness is determined by host-galaxy morphology, and also sets severe constraints on evolutionary schemes which attempt to link low-$z$ ULIRGs with RQQs.

Instead, we show that our results are as expected given the relationship between black-hole and spheroid mass established for nearby galaxies, and apply this relation to estimate the mass of the black hole in each object. The results agree remarkably well with completely-independent estimates based on nuclear emission-line widths; all the quasars in our sample have $M_{bh} > 5 \times 10^9 \, M_\odot$ while the radio-loud objects are confined to $M_{bh} > 10^9 \, M_\odot$. This apparent mass-threshold difference, which provides a natural explanation for why RQQs outnumber RLQs by a factor of 10, appears to reflect the existence of a minimum and maximum level of black-hole radio output which is a strong function of black-hole mass ($\propto M_{bh}^{-2.5}$). Finally, we use our results to estimate the fraction of massive spheroids/black-holes which produce quasar-level activity. This fraction is $\simeq 0.1\%$ at the present day, rising to $> 10\%$ at $z \simeq 2 - 3$.

Key words: galaxies: active – galaxies: photometry – infrared: galaxies – quasars: general – black hole physics

1 INTRODUCTION

Studies of the host galaxies of low-redshift quasars can enable us to define the subset of the present-day galaxy population which is capable of producing quasar-level nuclear activity. This is of obvious importance for constraining physical models of quasar evolution (Small \& Blandford 1992; Haehnelt \& Rees 1993; Kauffman \& Haehnelt 2000), and for exploring the connection between black-hole and galaxy formation (Silk \& Rees 1998, Fabian 1999, Franceschini et
al. 1999, Granato et al. 2001, Kormendy & Gebhardt 2001). Such observations are also of value for testing unified models of radio-loud AGN (e.g. Peacock 1987, Barthel 1989, Urry & Padovani 1995), constraining possible evolutionary links between ULIRGs and quasars (Sanders & Mirabel 1996), exploring the origin of radio-loudness (Blandford 2000), and as a means to estimate the masses of the central black holes which power the active nuclei (McLure et al. 1999).

Our view of low-redshift quasar hosts has been clarified enormously over the last five years, primarily due to the angular resolution and dynamic range offered by the Hubble Space Telescope (HST). After some initial confusion, recent HST-based studies have now reached agreement that the hosts of all luminous quasars ($M_V < -24.5$) are bright galaxies with $L > L^*$ (Bahcall et al. 1997, McLure et al. 1999, McLLeod & McLeod 2001). However, it can be argued, (with considerable justification) that this much had already been established from earlier ground-based studies (e.g. Smith et al. 1986, Véron-Cetty & Wolter 1990, Taylor et al. 1996).

In fact, as first convincingly demonstrated by Disney et al. (1995), the major advance offered by the HST for the study of quasar hosts is that it allows host galaxies to be mapped out over sufficient angular and dynamic range for a de Vaucouleurs $r^{1/4}$-law spheroidal component to be clearly distinguished from an exponential disc, at least for redshifts $z < 0.5$. This is not to suggest that AGN-host-galaxy morphological discrimination has proved impossible from the ground. Indeed for lower-luminosity AGN at $z < 0.1$, such as Seyfert galaxies, ground-based imaging has proved perfectly adequate for this task (e.g. Hunt et al. 1999) and in fact some early ground-based attempts to determine the morphology of low-redshift quasar hosts have also proved to be robust (e.g. Smith et al. 1986). However, to ensure an unbiased comparison of RQQ and RLQ hosts it is necessary to study host galaxies at $z > 0.15$ and to be able to determine host-galaxy morphologies for quasars with luminosities up to $M_V < -26$. Even by moving to the infrared to minimize nuclear:host ratio, Taylor et al. (1996) found that this could not be reliably achieved with typical ground-based seeing.

Nevertheless, great care needs to be taken to extract the full benefit of HST imaging of quasar hosts. In particular, deep observations are required to detect the extended low surface-brightness emission of even a massive host galaxy at $z \simeq 0.2$ to a radius of several arcsec from the nucleus. Unfortunately however, this inevitably leads to saturation of the nucleus, making accurate characterization of the luminosity of the central source impossible. This is crucial because, at the depths of interest for reliable host-galaxy characterization, scattered light in the WFPC2 PSF still makes a significant contribution to surface brightness out to an angular radius $\simeq 10$ arcsec (McLure, Dunlop & Kukula 2000). As demonstrated by McLeod & Rieke (1995), these problems of surface brightness bias, saturation, and inadequate knowledge of the large-angle properties of the true WFPC2 PSF, can explain much of the confusion produced by the first studies of quasar hosts undertaken after the correction of the HST optics with COSTAR (e.g. Bahcall, Kirkhados & Schneider 1994).

In this paper we present the final results from our 34-orbit Cycle-6 imaging study of quasar hosts, which was carefully designed to avoid these problems. Specifically, we acquired images of each quasar spanning a wide range of integration times (to allow an unsaturated, high dynamic-range image of each object to be constructed) and devoted an entire orbit to the construction of the necessary high dynamic-range PSF (via observations of a star of similar colour to the quasar nuclei, imaged at the same location on the same WF chip). Results from the first half of this programme were reported in McLure et al. (1999), where images for 19 objects from our 33-source sample were presented, modelled and analyzed. Here we present and model the images for the 14 targets which were observed in the latter half of 1998 and in 1999, and then summarize and discuss the results derived from the analysis of the completed sample. The results presented in this paper thus complete, extend and in several cases supersede those presented in McLure et al. (1999) (e.g. estimated black-hole masses for all objects are now calculated using more recent estimates of the black-holespheroid mass relation, yielding significantly lower values than were calculated by McLure et al. based on the relation presented by Magorrian et al. (1998)).

Several other substantial studies of low-redshift quasar hosts have now been undertaken with the HST (e.g. Bahcall, Kirkhados & Schneider 1997; Hooper, Impey & Foltz 1997; Boyce et al. 1998, McLeod & McLeod 2001). However, one unique feature of the present study is the deliberate focus on a comparison of the hosts of the three main classes of powerful AGN, namely radio-quiet quasars (RQQs), radio-loud quasars (RLQs) and radio galaxies (RGs). Moreover, we have ensured that this comparison can be performed in an unbiased manner by confining our sample to a narrow range in redshift ($0.1 < z < 0.25$) and requiring that the individual sub-samples are matched in terms of their luminosity distributions (optical luminosity in the case of the RQQ and RLQ sub-samples, and radio luminosity in the case of the RLQ and RG sub-samples - see Dunlop et al. (1993), McLure et al. (1999) and Section 2 for further details). Another strength of this study is the wealth of pre-existing data at other wavelengths, as detailed in Section 2. This has allowed us to maximise the impact of the HST imaging (through, for example, the determination of $R - K$ colours for all the host galaxies). Finally it is worth emphasizing that in this study we have sought to extract the properties of the stellar population which dominates the mass of the host galaxy. Thus, while we do include a statistical analysis of the prevalence of peculiar features such as tidal tails, we have endeavoured to minimize the distorting effect of the more transient activity by insisting on line-free imaging longward of the 4000Å break, and masking out obvious asymmetries prior to modelling the host morphology. This approach, coupled with the careful design of our observations, is the most likely explanation for why the results presented in this paper are generally cleaner, and more homogenous than the results of many other recent studies.

The layout of this paper is as follows. In Section 2 we review the main properties of the matched RG, RLQ and RQQ samples, and summarize the wealth of supporting ground-based multi-frequency data which now exists for these objects. In Section 3 we give details of the HST observations and briefly review the process of data reduction and PSF determination (McLure et al. 1999; McLure, Dunlop & Kukula 2000). Then, in Section 4 we give a brief description of the new VLA observations of the RQQs in our sample which es-
2 SAMPLE AND ASSOCIATED OBSERVATIONS

The HST imaging observations reported here complete the imaging of the full sample of 33 objects (10 RLQs, 13 RQQs and 10 RGs) defined for this study. This sample was selected from the slightly larger (40-source) sample imaged in the infrared by Dunlop et al. (1993) and Taylor et al. (1996) through the imposition of the slightly more restrictive redshift limits $0.1 < z < 0.25$. As described in McLure et al. (1999), this restriction in redshift range ensures that our R-band imaging through the F675W filter is not contaminated by the presence of strong emission lines such as [OIII], or Hα. The main effect of this additional redshift restriction is to exclude a small number of objects from the original sample of Dunlop et al. (1993) which have $0.25 < z < 0.35$. However, this has not significantly compromised the original statistical merits of this sample, namely that the RLQ and RQQ sub-samples are matched in terms of optical luminosity, and that the RLQ and RG samples are matched in terms of radio luminosity and radio spectral index (Dunlop et al. 1993).

Of the 33 sources in this sample, 19 were observed during the first year of our Cycle 6 allocation. The observations and analysis of these objects were presented by McLure et al. (1999). The remaining 14 objects for which the observations are presented and analyzed in this paper are listed in Table 1, along with the dates on which they were observed with the HST.

Also included in this paper are new, deep VLA observations of a subset of our RQQ sample. These 4.8 GHz observations, the results of which are presented in Section 4, go a factor of 3 deeper than the observations of Kukula et al. (1998) which were utilised by McLure et al. (1999). The important outcome of these observations is that we have now detected all but one of the RQQs in this sample at radio wavelengths. These new radio detections, coupled with completion of the HST observations, have allowed us to re-investigate and clarify a number of the relations between optical and radio properties which could only be tentatively explored by McLure et al. (1999).

We have also obtained improved infrared (UKIRT K-band) images for a small number of the more luminous quasars in our sample since the publication of McLure et al. (1999). These have been published in McLure, Dunlop & Kukula (2000), but the results of modelling these images are utilised in this paper to assist in the improved analysis of galaxy colours which is presented in Section 6.6.

Finally we note that Hβ emission-line spectroscopy of the RLQ and RQQ samples discussed here has recently been completed by McLure & Dunlop (2001). The results of this spectroscopic study are referred to in the discussion of black-hole mass estimation presented in Section 8.

3 HST OBSERVATIONS

The observations were made with the Wide Field & Planetary Camera 2 (WFPC2; Trager et al. 1994) on the Hubble Space Telescope through the F675W filter. The filter spans a wavelength range of 877Å from 6275.5 to 7152.5Å roughly equivalent to standard R band. This filter was selected in preference to a wider filter because it allowed both [OIII] and Hα emission lines to be excluded from the band-pass for source redshifts in the range $0.1 < z < 0.25$, and thus ensured that a clean measure could be made of the level of continuum light emitted by the quasar host galaxy at wavelengths longwards of the 4000Å break. As in McLure et al. (1999) target sources were centred on the WF2 chip, which was chosen in preference to WF1 or WF3 because of its marginally superior performance over the period immediately prior to our observations.

Deep sensitive images of the host galaxies are obviously

Table 1. Observing dates for the objects presented in this paper. Note that, despite its archive designation as 3C59, 0204+292 is in fact now classified as an RQQ (see Taylor et al. 1996, and the radio luminosity quoted in Table 3).

| Object          | HST Archive designation | Type | Observing date |
|-----------------|-------------------------|------|----------------|
| 0307+169        | 3C79                    | RG   | Jul 10 1998    |
| 0230−027        | PKS0230−027             | RG   | Sep 25 1998    |
| 1342−016        | PKS1342−016             | RG   | Dec 10 1998    |
| 1215+013        | PKS1215+013             | RG   | Jan 02 1999    |
| 1215−033        | PKS1215−033             | RG   | Jan 06 1999    |
| 1330+022        | PKS1330+022             | RG   | Apr 13 1999    |
| 1217+023        | PKS1217+023             | RLQ  | Jul 07 1998    |
| 1020−103        | PKS1020−103             | RLQ  | Jul 12 1998    |
| 2135−147        | PKS2135−14              | RLQ  | Oct 19 1998    |
| 2355−082        | PKS2355−082             | RLQ  | Oct 19 1998    |
| 0204+292        | 3C59                    | RQQ  | Jul 09 1998    |
| 1549+023        | 1E1549+023              | RQQ  | Sep 03 1998    |
| 2215−037        | EX2215−037              | RQQ  | Sep 26 1998    |
| 0052+251        | PG0052+251              | RQQ  | Nov 06 1998    |
Radio sources are detected within $\sim 1''$ of all bar one of the RQQs (positional accuracies are estimated to be within $\sim 100$ mas). Due to scheduling problems 1549+203 could only be allocated half the time given to the other five RQQs and the map of this object suffers from a correspondingly higher noise level.

5 NEW HST RESULTS

The WFPC2 F675W images, two-dimensional model fits, and the model-subtracted residual images of the 14 new objects are presented in Appendix A in Figs A1–A14, along with brief notes on each individual source. Comparable images and notes for the other 19 objects in the sample can be found in McLure et al. (1999). The observed luminosity profiles for the 14 new objects are presented in Appendix B in Figs B1–B14, along with the best-fitting model profiles extracted from the two-dimensional model fits.

Full details of the two-dimensional model procedure which we have used to determine the properties of the host galaxies can be found in McLure, Dunlop & Kukula (2000), along with the results of extensive tests of its ability to reclaim host-galaxy parameters from simulated data based on a wide range of host-galaxy:nucleus combinations at different redshifts. As emphasized in McLure, Dunlop & Kukula (2000), the success of this modelling depends on an accurate high dynamic range PSF, and the construction of an accurate error frame for each quasar image.

In brief, the modelling of the HST images was carried out in three separate stages. The first stage involves assessing how well the data can be reproduced assuming that the host galaxy is either an elliptical galaxy (with a surface-brightness distribution described by a de Vaucouleurs $r^{1/4}$-law) or a pure exponential disc. The remaining five parameters (host-galaxy position angle, host-galaxy axial ratio, host-galaxy scale-length, host-galaxy luminosity, and nuclear luminosity) are then adjusted until, when convolved with the PSF, the model best fits the data as determined by $\chi^2$-squared minimization (note that it is not assumed a priori that the radio galaxies have a negligible nuclear component). Then, if one assumed galaxy morphology yields a significantly better fit than the other, we can say that the galaxy is better described by a de Vaucouleurs law or by an exponential disc. As with all the modelling performed on the HST sample, once the minimum $\chi^2$ solution had been found, the modelling code was repeatedly re-started from close to the minimum $\chi^2$ solution, in order to ensure that the solution was stable.

The results of applying this procedure to the new HST images are given in Table 3, alongside the results already determined by McLure et al. (1999). The striking feature of these results, now confirmed with the complete sample, is that with the exception of 3 RQQs (0052+251 and the two lowest-luminosity RQQs 0257+024 and 2344+184) every single host galaxy is better described by a de Vaucouleurs law.

In our second approach we have removed the requirement of assuming that the host galaxy can be described as either a pure $r^{1/4}$-law or exponential disc, and allow a sixth parameter $\beta$ (where the luminosity profile of the galaxy is given by $I(r) \propto \exp(-r^\beta)$) to vary continuously. Thus $\beta = 1$
Table 2. High-sensitivity VLA observations of the six previously-undetected RQQs in our sample. Optical positions were measured from Digitised Sky Survey plates using the stsdas package in IRAF. All radio observations were made in C-band (4.8 GHz) with the VLA in A-configuration (angular resolution ≃ 0.4′′) on August 5 1999. The uncertainties in the measured flux densities are given as 3 times the rms noise in the image. Radio positions are accurate to within 100 mas.

| Quasar     | z    | $M_V$ | Optical Position (J2000) RA (h m s) Dec (° ′ ″) | Radio Position (J2000) RA (h m s) Dec (° ′ ″) | 4 GHz Flux Density /mJy | log($P_{4.8 GHz}$) |
|------------|------|-------|-----------------------------------------------|-----------------------------------------------|--------------------------|-------------------|
| 0244+194   | 0.176| −23.55| 02 47 40.85 +19 40 57.8                       | 02 47 40.84 +19 40 57.8                       | 0.18 ± 0.06              | 21.3              |
| PG 0923+201| 0.190| −24.56| 09 25 54.71 +19 54 04.4                       | 09 25 54.74 +19 54 05.0                       | 0.14 ± 0.06              | 21.3              |
| PG 0954+414| 0.239| −25.36| 09 56 52.35 +41 15 22.5                       | 09 56 52.39 +41 15 22.2                       | 0.25 ± 0.07              | 21.7              |
| 1549+203   | 0.250| −24.51| 15 52 02.36 +20 14 00.5                       | –                                              | −                         | <21.4             |
| 2215−037   | 0.241| −23.73| 22 17 47.77 −03 32 38.8                      | 22 17 47.72 −03 32 38.5                      | 0.13 ± 0.08              | 21.4              |
| 2344+184   | 0.138| −23.76| 23 47 25.71 +18 44 50.8                      | 23 47 25.77 +18 44 50.7                      | 0.19 ± 0.08              | 21.1              |

should result if the galaxy is best described by a pure exponential disc, and $\beta = 0.25$ should result if the galaxy really does follow a pure de Vaucouleurs law, but all values of $\beta$ are available to the program if this results in an improved quality of fit. The results of applying this procedure to the HST images are given in Table 4 (again for the complete sample) and illustrated in Fig 1. These results are discussed in more detail in Section 6.1.

Finally, an examination of Table 4 and Fig 1 reveals that whereas a very clean preference for $\beta = 0.25$ is displayed by the RGs and RLQ hosts, a few of the RQQ hosts (in particular, as mentioned above, 0052+251 and the two lowest-luminosity RQQs 0257+024 and 2344+184) have best-fit $\beta$ values which are intermediate between the values of 0.25 or 1.0 expected for pure elliptical or disc hosts. For this reason it was decided that the RQQ hosts should all be re-modelled with a 9 free-parameter fit, which allowed for the combination of both disc and bulge contributions to the host’s surface-brightness distribution. For nine out of the 13 RQQs this extra freedom still resulted in no significant disc component. However, for four objects (the above-mentioned three RQQs plus 0157+001) this procedure produced a significantly-improved model fit, and it is the $L_{host}$ and $L_{bulge}$ values from these combined fits which are used in all of the subsequent analyses. These four bulge-disc combinations are also noted in Table 3, and in that table (and in subsequent analysis) it is the scale-length, axial ratio and position angle of the dominant component which is adopted for these combined-fit objects.

The luminosity profiles, extracted from the two-dimensional model fits are presented for the 14 new objects in Appendix B. The profiles are followed out to a radius of $10''$, which is representative of the typical outer radii used in the modelling ($\langle r \rangle = 11''$).

6 ANALYSIS OF THE FULL SAMPLE

6.1 Host-Galaxy Morphologies

Examination of the results presented in Tables 3 and 4 confirms that the somewhat complex observing strategy outlined in Section 3 has successfully allowed the determination of host-galaxy morphology for all 33 objects in the sample. Also immediately apparent from Table 3 is that the huge amount of information available to the modelling code has not only allowed a clear morphological preference to be made, but can formally exclude the alternative host in all cases ($\Delta \chi^2 = 25.7$ corresponds to a 99.99% confidence level for a 5 parameter fit).

The results from the modelling of the RG and RLQ sub-samples are in good agreement with orientation-based unification, with all 20 objects found to have elliptical host galaxies. A perhaps more striking feature of these results is the extent to which the classic $r^{1/4}$ de Vaucouleurs law provides a near-perfect description of the host galaxies of the radio-loud objects (see Fig 1). As can be seen from Table 4 the best-fitting $\beta$ values for the combined RG and RLQ sub-samples all lie in the narrow range $0.19 < \beta < 0.26$, with the RG sample alone displaying an even narrower spread, $0.21 < \beta < 0.25$. This conclusion is further strengthened by a comparison of the beta histogram for the three sub-samples with that from the $\beta$-modelling tests performed by McLure, Dunlop & Kukula (2000). The application of the Kolmogorov-Smirnov test to the $\beta$ distribution of the 29 objects which were found to have single-component elliptical host galaxies, and those resulting from the $\beta$-model testing, returns a probability of $p = 0.23$. This confirms that, as far as the modelling code is concerned, any differences between these hosts and pure elliptical galaxies are not statistically significant. The best match between the test results and the actual data is for the RG sub-sample ($p = 0.39$), perhaps as would be expected considering the lack of a dominant point-source contribution.

Considering the long-standing belief that RQQs are often located in disc galaxies, the results from modelling of the RQQ sub-sample are quite unambiguous, with 9 of the 13 host galaxies showing no evidence for any disc component. Of the four objects which are best matched by a combined disc/bulge model, the luminosities of 0157+001 and 0052+251 are dominated by their bulge components, which respectively account for 83% and 71% of the total host luminosity. This means that the number of bulge-dominated RQQ hosts is 11 out of the 13 objects.

It is interesting to note that the two most disc-dominated galaxies, 0257+024 and 2344+184, are the hosts of by far the lowest luminosity AGN in the entire 33-object sample. In fact, converting their total luminosity (host+nucleus) from the model fits for these two objects to the equivalent absolute $V$-magnitudes gives $M_V = -22.6$ and $M_V = -21.9$ respectively (assuming $V - R = 0.8$ at
Table 3. The outcome of attempting to model the AGN host galaxies as either an exponential disc or a de Vaucouleurs spheroid. Source name and redshift are given in the first two columns, with the logarithm of radio luminosity given in column 3. The preferred host-galaxy morphology is given in column 4, with the $\Delta \chi^2$ between the chosen and alternative model listed in column 5. In column 6 $r_{1/2}$ is given irrespective of the chosen host morphology. Column 7 lists $\mu_{1/2}$ in units of $\text{mag arcsec}^{-2}$. Columns 8 and 9 list the integrated apparent magnitudes of the host galaxy and fitted nuclear component converted from F675W to Cousins $R$-band, while column 10 gives the ratio of integrated galaxy and nuclear luminosities. Columns 11 and 12 give the axial ratio and position angle (east of north) of the best-fitting host-galaxy model.

| Source | $z$ | log($P_{4\,GHz}$/WHz$^{-1}$sr$^{-1}$) | Host Morphology | $\Delta \chi^2$ | $r_{1/2}$/kpc | $\mu_{1/2}$ | $R_{\text{host}}$ | $R_{\text{nuc}}$ | L$_{\text{nuc}}$/L$_{\text{host}}$ | $b/a$ | PA/° |
|--------|----|-------------------------------|-----------------|----------------|-------------|-----------|--------------|-------------|-----------------|------|-------|
| RQQ    |     |                               |                 |                |             |           |              |              |                 |      |       |
| 0230−027 | 0.239 | 24.8 | Elliptical | 5900          | 7.7    | 21.8 | 17.5 | 0.95 | 113 |
| 0307−169 | 0.256 | 25.5 | Elliptical | 3500          | 9.4    | 21.4 | 17.2 | 20.9 | 0.03 | 1.00 | 13 |
| 0345+337 | 0.244 | 25.5 | Elliptical | 2400          | 13.1   | 23.3 | 18.0 | 21.1 | 0.06 | 0.70 | 90 |
| 0917+459 | 0.174 | 25.7 | Elliptical | 33000         | 21.9   | 23.0 | 16.1 | 19.4 | 0.05 | 0.76 | 36 |
| 0958+291 | 0.185 | 25.3 | Elliptical | 7800          | 8.5    | 22.0 | 17.1 | 18.5 | 0.27 | 0.95 | 40 |
| 1215−033 | 0.184 | 24.1 | Elliptical | 9300          | 8.5    | 22.0 | 17.1 | 22.3 | 0.008 | 0.87 | 60 |
| 1215+013 | 0.118 | 24.0 | Elliptical | 14000         | 4.7    | 21.0 | 16.5 | 19.9 | 0.05 | 0.94 | 142 |
| 1330+022 | 0.215 | 25.4 | Elliptical | 7400          | 15.7   | 22.9 | 17.1 | 19.5 | 0.11 | 0.79 | 79 |
| 1342−016 | 0.167 | 24.4 | Elliptical | 29000         | 23.8   | 22.9 | 15.6 | 21.8 | 0.003 | 0.93 | 96 |
| 2141+279 | 0.215 | 25.2 | Elliptical | 8500          | 24.8   | 23.5 | 16.7 | 25.6 | 0.0003 | 0.74 | 148 |
| RLQ    |     |                               |                 |                |             |           |              |              |                 |      |       |
| 0137+012 | 0.258 | 25.2 | Elliptical | 5100          | 14.2   | 22.6 | 17.2 | 17.3 | 0.8   | 0.85 | 35 |
| 0736+017 | 0.191 | 25.4 | Elliptical | 8900          | 13.3   | 22.9 | 16.9 | 16.2 | 1.9   | 0.97 | 13 |
| 1004−030 | 0.240 | 24.9 | Elliptical | 500           | 8.2    | 21.5 | 16.9 | 15.0 | 5.8   | 0.94 | 29 |
| 1020−013 | 0.197 | 24.7 | Elliptical | 4200          | 7.1    | 20.8 | 17.2 | 16.8 | 1.4   | 0.73 | 46 |
| 1217−023 | 0.240 | 25.9 | Elliptical | 2400          | 11.1   | 21.7 | 17.3 | 16.3 | 2.5   | 0.8  | 16 |
| 2135−147 | 0.200 | 25.3 | Elliptical | 2700          | 11.6   | 22.7 | 17.2 | 16.2 | 2.5   | 0.95 | 72 |
| 2141−175 | 0.213 | 24.8 | Elliptical | 570           | 8.2    | 21.2 | 17.3 | 15.9 | 3.7   | 0.47 | 118 |
| 2247+140 | 0.237 | 25.3 | Elliptical | 8100          | 13.5   | 22.4 | 17.2 | 16.9 | 1.3   | 0.63 | 118 |
| 2349−014 | 0.173 | 24.9 | Elliptical | 13000         | 19.2   | 22.7 | 15.9 | 16.0 | 0.9   | 0.89 | 45 |
| 2355−021 | 0.210 | 24.5 | Elliptical | 3000          | 10.4   | 22.0 | 17.1 | 17.4 | 0.77  | 0.73 | 177 |

$z = 0$, (Fukugita et al. 1995). Given that for this study the adopted quasar/Seyfert borderline is $M_V = -23.0$, it is clear that these two objects are not actually $b\text{ona fide}$ RQQs. The clear implication from this result is that all quasars, with $M_V < -23.0$, reside in luminous bulge-dominated hosts, irrespective of their radio power.

The morphological determinations for the RQQ host galaxies have placed on a firm footing the suggestion made previously by Taylor et al. (1996) and McLeod & Rieke (1995) that the probability of a RQQ having an early-type host was an increasing function of the quasar luminosity. Unlike the previous ground-based studies, the high resolution and temporally stable PSF offered by HST has permitted the confirmation of what were hitherto necessarily tentative conclusions, due to the uncertainties introduced by ground-based seeing conditions. Therefore, a strong conclusion from the morphological analysis of from the new HST images is that the radio luminosity of an AGN is not directly related to host-galaxy morphology. The relationship between host morphology and AGN luminosity is explored further in Section 7.

### 6.2 Host-Galaxy and AGN Luminosities

The host and nuclear luminosities are presented in the form of integrated absolute Cousins $R$-band magnitudes in Table 5, and in Figs 2 and 3. These have been derived from the apparent magnitudes listed in Table 3 which have been calculated by integrating the best-fit model components to infinite radius, and adopting the F675W flight-system zero-
**Table 4.** The outcome of the variable-\(\beta\) modelling. Column 2 lists the host morphology of the best fitting ‘fixed \(\beta\)’ model (results of which are given in Table 3). The best-fitting values for the \(\beta\) profile parameter are given in column 3. Column 4 gives \(\Delta \chi^2\), which quantifies the improvement in fit offered by the variable-\(\beta\) model over that already achieved with the best-fitting pure disc or elliptical model. Simulations indicate that \(\beta\) can be reclaimed to within a typical uncertainty of 0.01 – 0.02.

| Source | Host | \(\beta\) | \(\Delta \chi^2\) |
|--------|------|---------|------------------|
| RG     |      |         |                  |
| 0230–027 | Elliptical | 0.25 | 0                |
| 0307+169 | Elliptical | 0.21 | 52               |
| 0345+337 | Elliptical | 0.25 | 2                |
| 0917+459 | Elliptical | 0.23 | 264              |
| 0958+291 | Elliptical | 0.25 | 16               |
| 1215–033 | Elliptical | 0.24 | 0                |
| 1215+013 | Elliptical | 0.25 | 3                |
| 1330–022 | Elliptical | 0.24 | 16               |
| 1342–016 | Elliptical | 0.23 | 90               |
| 2141+279 | Elliptical | 0.25 | 2.2              |
| RLQ    |      |         |                  |
| 0137+012 | Elliptical | 0.19 | 126              |
| 0736+017 | Elliptical | 0.19 | 239              |
| 1004+130 | Elliptical | 0.25 | 5                |
| 1020–103 | Elliptical | 0.19 | 132              |
| 1217+023 | Elliptical | 0.26 | 2                |
| 2135–147 | Elliptical | 0.25 | 0                |
| 2141–175 | Elliptical | 0.28 | 23               |
| 2247+140 | Elliptical | 0.25 | 17               |
| 2349–014 | Elliptical | 0.26 | 10               |
| 2355–082 | Elliptical | 0.26 | 383              |
| RQQ    |      |         |                  |
| 0052+251 | Bulge/Disc | 1.09 | 267              |
| 0054+144 | Elliptical | 0.25 | 3                |
| 0157+001 | Bulge/Disc | 0.24 | 133              |
| 0204+292 | Elliptical | 0.24 | 216              |
| 0244+194 | Elliptical | 0.22 | 47               |
| 0257+024 | Disc/Bulge | 0.75 | 2792             |
| 0923+201 | Elliptical | 0.30 | 44               |
| 0953+415 | Elliptical | 0.27 | 9                |
| 1012–008 | Elliptical | 0.38 | 102              |
| 1549+203 | Elliptical | 0.25 | 0                |
| 1635+119 | Elliptical | 0.18 | 550              |
| 2215–037 | Elliptical | 0.25 | 0                |
| 2344+184 | Disc/Bulge | 0.43 | 1044             |

Figure 1. Sub-sample histograms of the best-fit \(\beta\) values from the variable-\(\beta\) modelling. The dotted line lies at \(\beta = 0.25\), corresponding to a perfect de Vaucouleurs law.

6.2.1 Host-Galaxy Luminosities

The mean and median integrated absolute magnitudes of the best-fit host galaxies in each sub-sample are:

- \(\langle M_\mathcal{R} \rangle = -23.53 \pm 0.09\) median\(=-23.52\) (ALL)
- \(\langle M_\mathcal{R} \rangle = -23.66 \pm 0.16\) median\(=-23.63\) (RG)
- \(\langle M_\mathcal{R} \rangle = -23.73 \pm 0.10\) median\(=-23.67\) (RLQ)
- \(\langle M_\mathcal{R} \rangle = -23.28 \pm 0.15\) median\(=-23.30\) (RQQ)

Two features of these results merit comment. First, the agreement between the absolute magnitudes of the host galaxies of the RG and RLQ sub-samples can be seen to be extremely good, with the median figures differing by only 0.04 magnitudes. This can be interpreted as strong evidence in favour of orientation-based radio-loud unification (see also Section 6.7). The second obvious feature of these results is that these new HST images appear to confirm the traditional finding that the hosts of RQQs are less luminous than those of RLQs, although the median difference of 0.37 magnitudes is a factor of two smaller than the difference typically claimed (e.g. Kirhakos et al. 1999). This does not seem to be

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an artifact of the inclusion of RQQ hosts with a substantial disc component, since if attention is confined to the 9 RQQs with solid elliptical model fits the values remain virtually unchanged, i.e.:

$$\langle M_R \rangle = -23.30 \pm 0.17 \quad \text{median} = -23.30 \quad \text{(RQQ)}$$

A large variation in the luminosity difference between RQQ and RLQ host galaxies has been reported in the literature. Differences have ranged from RQQ hosts being fainter than their RLQ counterparts by 0.7 → 1.0 magnitudes in optical studies (eg. Smith et al. (1986), Véron-Cetty & Woltjer (1990), Bahcall et al. (1997)), to no formal luminosity difference being detected from the near-infrared imaging of this HST sample (Taylor et al. 1996). It is therefore interesting to compare the difference in luminosity detected here with the results of the other two recent HST $R$-band imaging studies of Hooper et al. (1997) and Boyce et al. (1998), both of which also used two-dimensional modelling to analyse the host galaxies. Below we list the mean differences detected in the three studies ($\langle M_{RQQ} \rangle - \langle M_{RLQ} \rangle$) with the associated standard error. In calculating the figures for the Hooper et al. programme, where no attempt was made to distinguish the host galaxy morphologies, the luminosities of the best-fit $r^{1/4}$ model have been used.

$$\Delta M_R = 0.43 \pm 0.20 \quad \text{This Work}$$
$$\Delta M_R = 0.53 \pm 0.23 \quad \text{Hooper et al.}$$
$$\Delta M_R = 0.67 \pm 0.29 \quad \text{Boyce et al.}$$

### Table 5. Absolute magnitudes ($M_R$), and optical-infrared ($R - K$) colours of the best-fitting host galaxy and nuclear component for each AGN. Column 2 gives the $R$-band absolute magnitudes ($M_R$) of the total host galaxy derived from the current modelling of the HST data, assuming a spectral index of $\alpha = 1.5$ (where $f_\nu \propto \nu^{-\alpha}$). Column 3 gives the $R$-band absolute magnitudes ($M_R$) of the nuclear component as derived from the current modelling of the HST data assuming a spectral index of $\alpha = 0.2$. Columns 4 and 5 list the $R - K$ colours of the host galaxy and nuclear component respectively. These colours were derived by combining 12-arcsec aperture $R$-band photometry from our HST-based models with the 12-arcsec aperture $K$-band photometry derived by Taylor et al. (1996) and McLure, Dunlop & Kukula (2000), to minimize the uncertainty introduced by errors in constraining the galaxy scale-lengths at $K$ (see text for further details).

| Source  | $M_R$ (host) | $M_R$ (nuc) | $(R - K)$ $_{\text{Host}}$ | $(R - K)$ $_{\text{Nuc}}$ |
|---------|-------------|-------------|---------------------------|---------------------------|
| RG      |             |             |                           |                           |
| 0230–027| -23.55      | -           | 2.1                       | -                         |
| 0307+169| -23.96      | -19.94      | 2.1                       | 6.7                       |
| 0345+337| -23.05      | -19.63      | 3.6                       | 5.8                       |
| 0917+450| -24.20      | -20.65      | 3.4                       | 5.0                       |
| 0958+291| -23.36      | -21.70      | 1.9                       | 4.5                       |
| 1215+013| -22.85      | -19.35      | 2.5                       | 6.9                       |
| 1215–033| -23.29      | -17.87      | 2.6                       | 5.1                       |
| 1330+022| -23.70      | -21.00      | 2.8                       | 5.0                       |
| 1342–016| -24.58      | -18.35      | 2.6                       | 6.0                       |
| 2141+279| -24.09      | -14.94      | 3.3                       | 10.4                      |
| RLQ     |             |             |                           |                           |
| 0137+012| -24.04      | -23.53      | 2.8                       | 3.1                       |
| 0730+017| -23.58      | -24.01      | 3.2                       | 3.2                       |
| 1004+130| -24.10      | -25.70      | 3.0                       | 2.0                       |
| 1020–103| -23.36      | -23.47      | 2.3                       | 3.6                       |
| 1217+023| -23.71      | -24.40      | 3.2                       | 3.2                       |
| 2135–147| -23.37      | -24.09      | 2.5                       | 4.0                       |
| 2141+175| -23.43      | -24.52      | 2.7                       | 1.9                       |
| 2247+140| -23.80      | -23.80      | 2.8                       | 2.6                       |
| 2349–014| -24.32      | -24.00      | 2.9                       | 3.6                       |
| 2355–082| -23.62      | -23.06      | 2.7                       | 3.0                       |

| RQQ     |             |             |                           |                           |
| 0052+251| -23.33      | -24.39      | 2.5                       | 2.0                       |
| 0054+144| -23.61      | -24.49      | 3.1                       | 1.8                       |
| 0157+001| -24.52      | -23.70      | 2.8                       | 2.8                       |
| 0204+292| -23.30      | -23.03      | 2.5                       | 3.2                       |
| 0244+194| -22.77      | -23.24      | 2.3                       | 3.2                       |
| 0257+024| -23.46      | -19.69      | 2.6                       | 6.7                       |
| 0923+201| -23.25      | -24.57      | 3.2                       | 3.2                       |
| 0953+415| -22.86      | -25.52      | 3.0                       | 2.5                       |
| 1012+008| -23.78      | -23.95      | 3.2                       | 2.2                       |
| 1549+203| -22.26      | -23.98      | 3.4                       | 3.1                       |
| 1635+119| -23.05      | -21.54      | 3.3                       | 3.7                       |
| 2215–037| -23.52      | -22.58      | 2.7                       | 4.4                       |
| 2344+184| -22.97      | -20.30      | 2.6                       | 5.7                       |

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The clear implication from these results is that the host galaxies of RQQs are consistently fainter than those of RLQs by \( \sim 0.5 \) magnitudes in the \( R \)-band. Unlike in earlier studies, this difference can no longer be attributed to the model fitting of RQQs producing disc fits, since only 2 of the 26 RQQ host magnitudes included in the above figures are derived from an exponential host model. The common bias of the RLQ redshifts being consistently higher than those of the RQQs can also be firmly rejected as a possible cause of the difference, with the two quasar types having well-matched redshift distributions in all three studies. At least for the results presented here, and those of Hooper et al., the selection of the RLQ and RQQ quasar samples to have matched optical magnitudes, also excludes the possibility that the host magnitude difference is as a result of the RLQs being intrinsically more luminous. However, it is possible that the host-galaxy results of Boyce et al. may have been influenced by this effect, considering that the RLQs in that study are intrinsically 0.6 \pm 0.8 magnitudes brighter than the RQQs, although, as can be seen from the large error associated with this difference, the overlap of the total quasar luminosities is significantly larger than that of the hosts. It is worth noting at this point that although the results of Taylor et al. certainly do not formally support a difference in host magnitudes (with \( \Delta M_R = 0.4 \pm 0.3 \)), they are in fact perfectly consistent with the three sets of \( R \)-band HST results. The question of whether there is any detectable correlation between the host-galaxy and nuclear luminosities is explored below in Section 6.2.3.

A general result that can be taken from the luminosities of the best-fitting host galaxies is that, with the exception of the two Seyfert objects, all of them lie at the extreme end of the elliptical galaxy luminosity function. Adopting \( M_{R}^* = -22.2 \) (Lin et al. 1996), after having converted the published value of \( M_{K}^* = -21.8 \) to an integrated magnitude, it can be seen that all of the hosts have luminosities of \( L \geq L^* \), with 25 of the 33 having luminosities \( L \geq 2L^* \). These results are in good agreement with those of the previous \( K \)-band imaging study which also found all the hosts to be brighter than \( L_K^* \). Independent support for this result comes from the findings of Hooper et al. (1997) and Boyce et al. (1998), both of which used the F702W (wide R) filter on WFPC2. Using the Lin et al. value for \( M_K^* \), Hooper et al. found 15 of their 16 \( (z \approx 0.4) \) quasars to have \( L \geq L^* \), while Boyce et
al. found all 11 of their 14, objects for which a host model was fitted, to have $L > L^*$ (both sets of results having been converted to our adopted cosmology).

6.2.2 Nuclear Luminosities

The luminosities of the best-fitting nuclear components are listed in Table 5 and illustrated in Fig 3, with sub-sample averages summarized below. The radio galaxy 0230–027 has been excluded from these summary statistics as the modelling code detected no unresolved component in this object.

$\langle MR \rangle = -22.27 \pm 0.45$ median=$-23.36$ (ALL)

$\langle MR \rangle = -19.27 \pm 0.64$ median=$-19.64$ (RG)

$\langle MR \rangle = -24.07 \pm 0.22$ median=$-24.01$ (RLQ)

$\langle MR \rangle = -22.95 \pm 0.57$ median=$-23.70$ (RQQ)

As expected, the unresolved nuclear components displayed by the RGs are nearly two orders of magnitude weaker than those of the RLQs, in good agreement with the orientation-based radio-loud unification scheme first described by Peacock (1987) and Barthel (1989) (see Section 6.7). It would appear from the mean figures given above that there is a substantial difference in the nuclear components of the RQQ and RLQ quasar samples. However, the noticeable offset in the difference between the mean and median figures, combined with examination of histogram shown in Fig 3, reveals that the RQQ mean is being biased by the three objects which have total luminosities fainter than $M_V = -23.0$. If these three objects are excluded from the RQQ sample, the mean and median values of nuclear luminosity become:

$\langle MR \rangle = -23.95 \pm 0.26$ median=$-23.97$ (RQQ)

completely consistent with the equivalent RLQ figures. The similarity between the nuclear components of the RQQ and RLQ sub-samples is reassuring considering that the sample-selection process was indeed originally designed to produce two quasar samples well-matched in luminosity. This result also re-affirms that the measured difference in the respective host magnitudes of the RQQs and RLQs cannot be attributed to some bias arising from the RLQs in our sample having, on average, more powerful nuclei.

6.2.3 Quasar Host-Nuclear Luminosity Correlation

Because the measured difference between the RQQ and RLQ host-galaxy magnitudes is substantially greater than the average difference between their respective nuclear components (excluding the three sub-luminous objects) it is not expected that there should be a strong host-nuclear luminosity correlation. This is indeed the case, and application of the Spearman Rank correlation test confirms that there is no evidence for any correlation, returning a probability $p = 0.95$ that the null hypothesis of no correlation is acceptable.

The null result found here contrasts with the positive correlations found previously by many authors (eg. Smith et al. (1986), Bahcall et al. (1997), Hooper et al. (1997)). The suspicion often cast upon the reality of a positive host-galaxy:nuclear correlation is that there are two obvious selection effects at work. The first of these is that it is obviously much more difficult to detect faint galaxies which host brighter quasars. It seems clear that given the difficulties associated with ground-based seeing, and the perils of PSF-subtraction on HST data (eg. Bahcall et al. 1997), that this could well be a contributing factor. In the present study there are several RQQs with bright nuclear components, and relatively faint host galaxies (eg. 0953+415), providing confidence that the techniques employed here have been successful in overcoming this possible source of a false correlation. The other selection effect which could contribute to a false positive result is that weak AGN in bright host galaxies will not be classified as quasars.

Although the results presented here for powerful low-$z$ quasars show no evidence for a correlation, it does appear to exist for lower-luminosity Seyfert galaxies. In the case of Seyferts, the reduction in both redshift and contrast between the host and nuclear components means that it can be confidently assumed that all of the host galaxies are being detected, removing one of the sources of bias mentioned above. Indeed, the work of McLeod & Rieke (1994) on two samples of low- and high-power AGN at low-z has led them to suggest that the relation between host and nuclear luminosity is of the form that there is a minimum host luminosity required to produce a particular quasar luminosity. This would be consistent with both the positive correlation found at lower AGN power, and the flat relation found here for higher powered quasars. The host galaxies for the quasars studied here have already been shown to be among the brightest known, and therefore demand that the galaxy-quasar relation tails-off at high quasar luminosity. The question of how the nuclear and host-galaxy luminosities relate to central black-hole mass and quasar fuelling rate is pursued in Section 8.

6.3 Scale-lengths

The best-fitting values of the host-galaxy scale-lengths listed in Table 3 are displayed as histograms for the separate sub-samples in Fig 4. The average figures for the three sub-samples are:

$\langle r_{1/2} \rangle = 12.23 \pm 1.00$ median=$10.45$ (ALL)

$\langle r_{1/2} \rangle = 13.76 \pm 2.18$ median=$11.27$ (RG)

$\langle r_{1/2} \rangle = 11.73 \pm 1.07$ median=$11.28$ (RLQ)

$\langle r_{1/2} \rangle = 11.45 \pm 1.66$ median=$9.30$ (RQQ)

Taking the median figures, the RG and RLQ sub-samples can again be seen to be in remarkable agreement, while the RQQ hosts are $\sim 20\%$ smaller.

Galaxy half-light radii of $\sim 10$ kpc mark out these radio galaxies and quasar hosts as being giant elliptical galaxies. In a study of the galaxies in the Virgo cluster, Capaccioli et al. (1992) found that beyond a scale-length of $\sim 3$ kpc the only galaxies to be found were extremely luminous ellipticals, or cD type brightest cluster galaxies (BCG). The question of how the properties of the host galaxies studied here relate to those of BCGs is explored further in Section 7. In further support of the radio-loud unification scheme the scale-length figures obtained here for the RG and RLQ sub-samples can be compared with the B- and V-band study of Smith & Heckman (1990), who determined the scale-lengths of 41 powerful radio galaxies in the redshift range $0 < z < 0.26$. The median scale-length figure determined is equivalent to $\sim 17$ kpc in the cosmology adopted here, somewhat larger
than the results presented above. However, if the comparison is restricted to the 22 objects studied by Smith & Heckman which have strong optical emission lines, more likely to have the same FR II morphology as the majority of radio-loud objects studied here, then the median scale-length falls to 13 kpc, in excellent agreement with the figures presented above.

If the host scale-lengths are plotted against the integrated model luminosities, as in Fig 5, a tight correlation is found, and a least-squares fit to these data produces a relation of the form $L \propto r^{0.75}$. This is in impressively close agreement with the corresponding result $L \propto r^{0.70}$ found for low-$z$ inactive ellipticals by Kormendy (1977). Given this result, it is obviously of interest to ask whether the best-fit model parameters for $\mu_{1/2}$ and $r_{1/2}$ given in Table 3 do in fact yield a Kormendy relation of slope $\sim 3$.

6.4 The Kormendy Relation

The $\mu_{1/2} - r_{1/2}$ relation for the host galaxies of all 33 objects is shown in Fig 6. For the four RQQ sources which have combined disc/bulge fits it is the parameters for the bulge component that have been plotted. The least-squares fit to the data (solid line) has the form:

$$\mu_{1/2} = 2.90_{-0.22}^{+0.22} \log_{10} r_{1/2} + 18.35_{-0.22}^{+0.22}$$ (ALL)

(errors are $\pm 1\sigma$) showing for the first time from optical imaging, that the host galaxies of quasars lie on this photometric projection of the fundamental plane. The individual fits to the $\mu_{1/2} - r_{1/2}$ relations for the individual sub-samples are:

$$\mu_{1/2} = 2.86_{-0.32}^{+0.32} \log_{10} r_{1/2} + 18.44_{-0.36}^{+0.36}$$ (RG)

$$\mu_{1/2} = 3.98_{-0.71}^{+0.71} \log_{10} r_{1/2} + 17.02_{-0.75}^{+0.75}$$ (RLQ)

$$\mu_{1/2} = 2.99_{-0.34}^{+0.34} \log_{10} r_{1/2} + 18.39_{-0.30}^{+0.30}$$ (RQQ)

It is readily apparent from the individual Kormendy relations that the RG and RQQ sub-samples follow very similar relations, consistent with each other in terms of slope and normalization. One encouraging aspect of this is that the best-fitting bulge components for the RQQ objects which have combined disc/bulge model fits lie naturally on the Kormendy relation defined by the other nine, single component, RQQ hosts. With no restrictions placed on the range of parameter values available to the modelling code, the fact that the fitted bulge components have physically sensible values gives further confidence that the introduction of the combined disc/bulge fits was justified, and does not involve over-fitting of the data.

It would appear from the individual fits given above that the RLQ sub-sample follows a steeper slope, compared with the best-fit relation to the other two sub-samples. It is also noticeable that the $1\sigma$ error returned by the least-squares fitting procedure is more than twice that returned from the fitting to the RG sub-sample (which also has 10
modelling follows simply from the fact that the integrated luminosity of a galaxy model follows the relation:

\[ L_{\text{int}} \propto I_{1/2} r_{1/2}^2 \]  

(2)

where \( I_{1/2} \) is the surface-brightness at \( r_{1/2} \). Given that \( \mu_{1/2} \propto -2.5 \log I_{1/2} \), it follows that if the modelling procedure can successfully constrain the host luminosity, but not the scale-length, then the apparent best-fitting values of \( \mu_{1/2} \) and \( r_{1/2} \) will be randomly distributed along a relation obeying:

\[ \mu_{1/2} \propto 5 \log r_{1/2} \]  

(3)

with appropriate normalization to fit the integrated luminosity. Due to the fact that the host galaxies display a relatively small spread in luminosity (mean \( M_R = -23.44, \sigma = 0.64 \)) it might be expected that, if our analysis was unable to accurately determine the individual host scale-lengths, the resulting \( \mu_{1/2} - r_{1/2} \) relation would be well fitted with a slope of 5 (as found by de Vries et al. 2000) and a normalization to match \( M_R = -23.44 \). Such a relation has the form:

\[ \mu_{1/2} = 5.0 \log_{10} r_{1/2} + 16.40 \]  

(4)

and is plotted as the dotted line in Fig 6. It is obvious from Fig 6 that this relation is not consistent with the data, giving confidence that the methods of analysis employed here have allowed the accurate determination of the host-galaxy scale-lengths.

The diagnostic power of Fig 6 is discussed in the context of recent studies of the fundamental plane in Section 7.2.

### 6.5 Axial Ratios

If the AGN host galaxies are indeed indistinguishable from massive inactive ellipticals then they should display an axial ratio distribution which is also identical. The axial ratio distribution of normal ellipticals is well studied and known to peak at values of \( b/a \geq 0.8 \) (Sandage, Freeman & Stokes 1970; Ryden 1992). The best-fitting axial ratios from the two-dimensional modelling are listed in Table 3 and displayed in the form of separate sub-sample histograms in Fig 7. The corresponding mean and median values are:

- \( b/a = 0.81 \pm 0.02 \) median=0.85 (ALL)
- \( b/a = 0.86 \pm 0.03 \) median=0.90 (RG)
- \( b/a = 0.80 \pm 0.05 \) median=0.82 (RLQ)
- \( b/a = 0.78 \pm 0.03 \) median=0.84 (RQQ)

It can be seen from these figures that the axial ratio distributions of all three sub-samples are perfectly consistent with that expected from a sample of elliptical galaxies. There is a slight suggestion that the RGs have higher axial ratios than average, but this is not formally significant.

The axial ratio results from the two-dimensional modelling agree well with the recently-published findings of Boyce et al. (1998) who found that all 11 of the quasars from their 14-object \((z \approx 0.3)\) sample for which a model fit was possible, displayed axial ratios with \( b/a > 0.65 \). In contrast, Hooper et al. (1997) found only 2 of their 16 \((z \sim 0.45)\) quasars to have axial ratios with \( b/a > 0.6 \). The reasons for this apparently contradictory result probably lie in a combination of the higher redshifts of the Hooper et al. objects,
Figure 7. The axial ratio distributions for the three host-galaxy sub-samples as determined by the two-dimensional modelling.

Figure 8. The rest-frame $R-K$ colours for the three host-galaxy sub-samples. The three sub-samples can be seen to be consistent with each other, tightly distributed around a value of $R-K \approx 2.5$.

and their use of the PC instead of the WF detectors. As has been explained, these two factors will undoubtedly result in a reduced sensitivity to low surface-brightness features, given that the exposure times used were the same as adopted in this study. In fact, Hooper et al. noted their concern that their modelling technique of two-dimensional cross-correlation could have been influenced by high surface-brightness features such as bars or tidal tails. Given the clear result presented here, and the support of the Boyce et al. results at similar redshifts, this seems the most likely explanation of the Hooper et al. result.

6.6 Host-Galaxy Colours

The results presented in the last five sub-sections strongly suggest that in terms of morphology, luminosity, scale-length, Kormendy relation, and axial-ratio distribution the hosts of powerful AGN are identical to normal, massive, inactive ellipticals. The one final parameter which can be readily recovered from the modelling is host-galaxy colour. The desire to obtain reliable optical-infrared colours for the host galaxies was one of the original motivations for this HST imaging study. Given that the results presented thus far suggest the host galaxies are otherwise normal massive ellipticals, it is now even more interesting to investigate whether the hosts also display the red colours associated with old stellar populations, or whether they have significantly bluer colours, indicative of either a generally young stellar population, or of substantial secondary star formation induced by interactions, the central AGN, or both.

The simplest and most obvious way to calculate $R-K$ colours for the host galaxies is to combine the integrated optical magnitudes presented in Table 5 with the integrated absolute $K$-band magnitudes derived for the sample by Taylor et al. (1996). However, there are two potentially serious problems with this strategy. First, Taylor et al. did not always decide on the same host morphology as has now been revealed by the HST imaging, and so in some cases their quoted absolute $K$ magnitude is derived from, for example, a disc fit which originally appeared marginally preferable to a de Vaucouleurs law. This obviously matters because the disc fits to the $K$-band data were $0.5 \rightarrow 1.0$ magnitudes fainter than the equivalent elliptical fits. Second, in
the light of the new, better constrained models derived from the HST data, and the results of more recent tip-tilt infrared imaging reported by McLure, Dunlop & Kukula (2000), it is now clear that the half-light radii derived by Taylor et al. were systematically over-estimated (most likely due to the \(\sim 1''\) seeing and coarse spatial resolution (0.62''/pix) of the original K-band observations - see also Simpson et al. 2000). Given that the integrated luminosity of the standard de Vaucouleurs \(r^{1/4}\) law is proportional to \(I_{\lambda1/2r^2}\), it is clear that scale-length errors of the order of \(1.5 - 2\) could seriously bias the integrated K-band luminosity.

Therefore, to obtain a robust value for the \(R - K\) colour of each host from the existing data we proceeded as follows. First, the K-band luminosities were based on the best-fitting model with the same morphology as the equivalent HST result. Second, the \(R - K\) colours were based on 12''-diameter aperture photometry performed on both the HST and K-band model fits, to minimize the impact of scale-length errors on the derived colour, while at the same time including the vast majority (> 75%) of host-galaxy light.

The results of these host-galaxy \(R - K\) calculations are listed in Table 5, with absolute \(R - K\) histograms for the separate sub-samples shown in Fig 8. The absolute colours have been calculated assuming spectral indices of \(\alpha = 1.5\) & \(\alpha = 0.0\) for the \(R\) and \(K\)-bands respectively.

It is readily apparent from Fig 8 that the colour distributions of the three sub-samples are consistent with each other, with mean and median absolute colours:

\[
\begin{align*}
(R - K) &= 2.48 \pm 0.05 & \text{median} &= 2.44 & \text{(All)} \\
(R - K) &= 2.47 \pm 0.09 & \text{median} &= 2.56 & \text{(RG)} \\
(R - K) &= 2.48 \pm 0.07 & \text{median} &= 2.46 & \text{(RLQ)} \\
(R - K) &= 2.48 \pm 0.08 & \text{median} &= 2.41 & \text{(RQQ)}
\end{align*}
\]

As demonstrated by Nolan et al. (2001) such colours are consistent with those expected of an evolved stellar population of age 10 – 13 Gyr. The homogeneity of these colours provides a further strong indication that the hosts of all three classes of AGN are derived from the same parent population of massive, well-evolved elliptical galaxies. The inevitable conclusion from this is that these galaxies, or at least the bulk of their stellar populations, must have formed at high redshift, a theme which is revisited in Section 9. The fact that the \(R - K\) colours are so similar to passive ellipticals forces the further conclusion that any star-formation associated with AGN activity must either be dust enshrouded, confined to tidal features masked from the modelling, or possibly restricted to the central \(\sim 1\) kpc of the host (in which case the additional emission would almost certainly be attributed to the unresolved nuclear component). Large-scale star-formation involving a substantial fraction of the mass of the host galaxy, is effectively ruled out by these results.

6.7 Nuclear Colours: Implications for Unified Models of Radio-Loud AGN

The \(R - K\) colours of the fitted nuclei for all 33 objects (except the RG 0230–027 - see above) are plotted against absolute nuclear \(R\) magnitude in Fig 9. It is clear from this plot that there is no significant difference between the colours of the nuclei of the RQQs or RLQs. More importantly, however, this plot allows us to explore the relationship between the relatively weak nuclear components found in the RGs, and the (naturally) brighter components found in the RLQs.

This is of value because it allows us to perform another test of the proposed unification of RLQs and RGs by orientation (Peacock 1987, Barthel 1989). As discussed above, our host-galaxy results can already be viewed as providing strong support for such a picture, due to the fact that the basic properties (i.e. luminosity, size, axial ratio, age etc) of the RGs and RLQ hosts in our matched samples are indistinguishable. However, it must also be realised that identical RG and RLQ host galaxy properties would also be expected in a scenario where the two classes of radio-loud AGN are in fact linked by evolution in time, rather than by orientation.

Fig 9 offers the possibility of differentiating between these two alternative explanations, by allowing us to test whether the \(R - K\) colours of the RG nuclear components are consistent with dust-obscured quasar nuclei. If it is assumed that the reddening of the RG nuclei obeys a \(\kappa_{\lambda} \propto \lambda^{-1}\) law, then the reddening vector which should connect the RGs and RLQs in Fig 9 would be expected to be given by \(E(R - K) = 0.70A_R\).

In fact the slope of the least-squares fit in Fig 9 (solid line) is \(0.68 \pm 0.05\), which corresponds to a dust reddening law \(\kappa_{\lambda} \propto \lambda^{-0.95\pm0.15}\). Furthermore, if the fitted \(R\)-band nuclear components of the RGs are de-reddened using a slope of 0.68, until each object has the mean colour of the RLQ nuclei \((R - K = 3.0)\), then the mean luminosity of the de-reddened RG nuclei is \(M_R = -23.90 \pm 0.46\). It can be seen that this is perfectly consistent with the previously determined mean nuclear luminosity of the RLQ sample, \(M_R = -23.73 \pm 0.10\). Given the clean nature of this result, there seems little room.
for argument that the RG and RLQ sub-samples are most consistent with being drawn from the same population of objects, viewed from different orientations, with every RG displaying evidence of a buried quasar nucleus.

7 THE RELATION OF QUASAR HOSTS TO ‘NORMAL’ GALAXIES

7.1 Dependence of Host Morphology on Nuclear Luminosity

In the preceding section we presented a series of results which together lead to the inescapable conclusion that the hosts of virtually all powerful AGN are essentially normal, massive ellipticals. This result is particularly clean for the radio-loud objects in our sample. However, exceptions to any rule are often of great significance, and one should not ignore the fact that we have found two disc-dominated hosts in our RQQ sample, along with two bulge-dominated RQQ hosts which do, however, possess significant disc components.

What might such exceptions to the general ‘perfect elliptical’ rule be telling us? An important clue comes from the fact that, as discussed above, the two disc-dominated hosts in our sample transpire to contain the two least-luminous nuclei. Indeed, in much of the preceding analysis, these two objects have had to be deliberately excluded from sample comparisons due to the fact that they are not actually luminous enough to be classed as quasars. However, the fact that these two radio-quiet low-luminosity interlopers are the only objects have had to be deliberately excluded from sample comparisons due to the fact that they are not actually luminous enough to be classed as quasars. However, the fact that these two radio-quiet low-luminosity interlopers are the only RQQs in our sample with disc-dominated hosts strongly suggests that host morphology is a function of nuclear optical luminosity.

Our own samples do not span a sufficient range in optical luminosity to allow a statistical test of this hypothesis. However, we can explore this possibility, and place our results in a wider context, by combining them with those of Schade et al. (2000) who have studied the morphologies of the host galaxies of a large sample of X-ray selected AGN spanning a wider but lower range of optical luminosities. This we have done in Fig 10, which shows the ratio of bulge-to-total host luminosity plotted as a function of nuclear optical power.

This instructive diagram illustrates a number of important points. First, it shows that while pure spheroidal galaxies can host AGN with optical luminosities ranging over two orders of magnitude, very disc-dominated hosts start to die out for $M_V(\text{nuc}) < -21$ and no disc-dominated galaxies appear capable of hosting nuclear emission more luminous than $M_V \approx -23$ (close to the traditional, albeit rather ad hoc, quasar:seyfert classification boundary). Second, this plot confirms that the two most ‘discy’ hosts uncovered in the present study do indeed lie in a region of the parameter space where disc-dominated host galaxies are quite common. Third, in the context of this diagram, the persistence of significant discs around two of the more luminous RQQs in our sample (0052+251 and 0157+001) seems perfectly natural, and certainly not inconsistent with the clear trend towards more spheroid-dominated hosts with increasing nuclear output.

Fig 10 clarifies the origin of much of the confusion which has surrounded the nature of RQQ hosts. In particular it is clear that studies of RQQ samples which extend to include significant numbers of objects with nuclei less luminous than $M_V \approx -23$ are likely to uncover significant numbers of hosts with substantial discs, in contrast to the results of the present study. This does indeed appear to be the case. For example, Bahcall et al. (1997) and Hamilton et al. (2002) have reported that approximately one third to one half of radio-quiet quasars lie in disc-dominated hosts. However, if one restricts their results to quasars with nuclear magnitudes $M_V < -23.5$ (for which we find 10 out of 11 RQQs lie in ellipticals), it transpires that 6 out of the 7 remaining objects in the Bahcall et al. study lie in ellipticals, while 17 out of the remaining 20 objects in the archival study of Hamilton et al. also have elliptical host galaxies.

Fig 10 is also at least qualitatively as expected if AGN are in fact drawn at random from the general galaxy population, subject to the constraint that the host galaxy contains a black hole of sufficient mass to produce the observed nuclear output (see also Wisotzki, Kuhlbrodt & Jahnke 2001). As discussed further in Section 8, a black-hole mass $m_{bh} > 5 \times 10^8 M_\odot$ appears to be required to produce a quasar with $M_V < -23.5$. When this constraint is combined with the now apparently inescapable result that black-hole mass is proportional to spheroid mass $M_{bh} \simeq 0.001 - 0.003 M_{sph}$, then Fig 10 can be viewed as a simple manifestation of the fact that spheroids in the present-day universe with baryonic...
masses $> 3 \times 10^{11} M_\odot$ are rarely accompanied by significant discs (and are thus classed as massive ellipticals).

This conclusion leads naturally to the question of why most massive black holes in the present-day universe should be inactive, while a subset are receiving sufficient fuel to shine as quasars. Below we present the results of a series of statistical comparisons designed to decide whether there are in fact any significant observable differences between the hosts (and host environments) of quasars and comparably massive, but inactive, elliptical galaxies.

### 7.2 Comparison with ULIRG merger remnants

Before proceeding to compare quasar hosts with old passively-evolving massive ellipticals, it is worth pausing to consider whether host galaxies with the properties we observe could plausibly be produced as the outcome of a (relatively recent) violent merger of two disc galaxies. This might already seem fairly implausible given the series of results presented in Section 6, in particular the red host-galaxy colours consistent with old evolved populations. Nevertheless, it is well known from both simulations (Barnes & Hernquist 1992) and observations that the red→infrared light from ULIRGs often follows an $r^{1/4}$ law (Wright et al. 1990, Scoville et al. 2000) and it is possible that dust extinction in such objects could also produce red colours (albeit such colours would be unlikely to mimic so well the properties of an evolved population, as discussed by Nolan et al. 2000). Moreover, several authors have proposed an evolutionary scheme in which ULIRGs might evolve into radio-quiet quasars on a time-scale $< 1$ Gyr (Sanders & Mirabel 1996).

Fortunately, the data on quasar hosts presented here are of sufficiently high quality to completely exclude the possibility that they are the recent outcome of the violent disc-disc mergers which appear to be the origin of most ULIRGs in the low-redshift universe. While it is true that ULIRGs such as Arp220 have surface brightness profiles well-described by an $r^{1/4}$-law, recent observations (Genzel et al. 2001) have shown that such remnants lie in a completely different region of the fundamental plane than that which we have shown is occupied by quasar hosts. Specifically, the effective radii of ULIRGs is typically $\simeq 1$ kpc, an order of magnitude smaller than the typical effective radii of the quasar hosts as illustrated in Figs 4, 5 and 6.

Indeed, one can go further and conclude that whereas ULIRGs can be the progenitors of the population of intermediate-mass ellipticals which display compact cores and cusps (Faber et al. 1997), the quasar hosts lie in a region of the $\mu_e - r_e$ plane which is occupied by boxy ellipticals with large cores, a large fraction of which lie in the centres of clusters (Faber et al. 1997). This is demonstrated in Fig 11 where we have constructed a composite $R$-band Kormendy diagram, combining our own results on the 31 bulge-dominated AGN hosts with the data from Genzel et al. (2001) on LIRGs and ULIRGs, and the data from Faber et al. (1997) on ‘discy’ and ‘boxy’ ellipticals.

This plot provides strong evidence that the quasar hosts belong to the class of large ‘boxy’ ellipticals, which various authors have suggested form at earlier times than their lower-mass ‘discy’ counterparts (Kormendy & Bender 1996, Faber et al. 1997).

In summary, all the available evidence indicates that luminous low-redshift quasars are the result of the re-triggering of a massive black hole at the centre of an old evolved elliptical, and do not generally occur in the remnants of the recent disc-disc mergers which power ULIRGs. While there is one RQQ in our sample which is also a ULIRG (0157+001), in this case it is clear that at least one of the merging galaxies already has a massive spheroid capable of containing a pre-existing massive black hole. Thus, while there is clearly some overlap between the ULIRG and quasar phenomenon, present evidence suggests that, at least at low redshift, the ULIRG → quasar evolutionary route can only apply to a fairly small subset of objects.

### 7.3 Comparison with Brightest Cluster Galaxies

Having established that the hosts of quasars are massive ellipticals, it is then of interest to explore how their properties compare to those of inactive comparably-massive ellipticals at similar redshift. In this sub-section we compare the properties of the quasar hosts, and in particular their scale-lengths, with those of bright galaxies found in rich clusters. This comparison is motivated, in part, by the location of the host galaxies on the $\mu_e - r_e$ projection of the fundamental plane as discussed in the previous subsection. However, this motivation is further strengthened by the results of a study of the environments of the quasars and radio galaxies undertaken by McLure & Dunlop (2000). McLure & Dunlop

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Figure 11. A comparison of the properties of the AGN hosts with those displayed by various other types of spheroid on the photometric projection of the fundamental plane. Symbols for the quasar hosts and radio galaxies are as in Figs 5 and 6. The stars are the data for ULIRGs and LIRGs from Genzel et al. (2001) transformed from the infrared to the $R$-band assuming $R - K = 2.5$. Triangles and squares indicate the positions of ‘discy’ and ‘boxy’ ellipticals from Faber et al. (1997) after conversion to $H_0 = 50$ km s$^{-1}$ Mpc$^{-1}$.
(2000) used the full WFPC2 images of our AGN sample to measure the spatial clustering amplitude \( B_{qg} \). This quantity is inevitably rather poorly constrained for each individual object, but, on average, they found that all three classes of AGN typically inhabit environments as rich as Abell Class \( \approx 0 \). Moreover, several objects appeared to lie in noticeably richer environments (Abell Class 1).

Given these clustering results, and the fact that the quasar hosts are clearly comparable to many BCGs in terms of optical luminosity, it is interesting to explore whether the sizes of the quasar hosts are as expected in the light of their apparent cluster environment. Strong circumstantial evidence that the quasar host galaxies are directly comparable to first-ranked cluster galaxies comes from a comparison of the host-galaxy Kormendy relation with that found for cD galaxies by Hamabe & Kormendy (1987) in the B-band:

\[
\mu_{1/2} = 2.94 \log_{10} r_{1/2} + 20.75
\]

If a typical elliptical galaxy colour at \( z = 0.2 \) of \( B - R = 2.4 \) is assumed (Fukugita et al. 1995), then the inferred B-band Kormendy relation for all 33 objects in the HST sample presented in Section 6.4 becomes:

\[
\mu_{1/2} = 2.90 \log_{10} r_{1/2} + 20.75
\]

which can be seen to be in excellent agreement with the Hamabe & Kormendy result.

To explore further the connection between host cluster environment and galaxy size, a sample of 51 first-ranked cluster galaxies (confined to the same redshift range as the HST sample but spanning a range of Abell classes) was constructed from the study of Hoessel & Schneider (1985). Application of the KS test to the scale-length distributions of the AGN host galaxies, and the 14 first-ranked cluster galaxies in this sample from Abell clusters of class 0 & 1, shows that the two are not significantly different \((p = 0.22)\). However, the extension of the comparison sample to include the first-ranked galaxies of Abell class 2 clusters (15 objects), shows these two distributions to be different at the 2\( \sigma \) level \((p = 0.011)\).

This is an very interesting result because it provides independent support for the results of the cluster analysis of McLure & Dunlop (2000) which also favour an environment more consistent with Abell class 0, or at most (in only a few cases) Abell class 1.

Finally, to check that the ground-based results discussed above can be confirmed by modelling of more recent HST data, and to minimize possible systematic errors, we decided to model a series of HST images of bright cluster galaxies (at comparable redshift to our quasars) using the same 2-D modelling code used to model the AGN hosts. HST WFPC2 images of five Abell clusters with similar redshifts to the objects in the HST sample were therefore retrieved from the HST archive facility\( ^{\dagger} \). Following the production of reduced images of comparable signal-to-noise to the host galaxy images, the four brightest galaxies in each cluster were identified and modelled in identical fashion to the AGN hosts. The results of the elliptical model fits to these 20 galaxies are presented in Table 6.

BCGs in Table 6 can be instructively compared with the corresponding quantities for the AGN hosts listed in Tables 3 and 5. First, it is clear that none of our AGN hosts is comparable in size to either the first or second ranked cluster galaxies in very rich, Abell Class 4 clusters. However, the largest of the AGN hosts are comparable in size to the 1st or 2nd ranked galaxies in the Abell Class 0 and Class 1 clusters. This more detailed galaxy-by-galaxy comparison thus provides further support for the basic conclusions of McLure & Dunlop (2000) already mentioned above. Indeed, at least within the radio-loud subsamples, there is a suggestion of a rather clean link between the biggest galaxies and the richest cluster environments. Specifically, the two radio galaxies found by McLure & Dunlop (2000) to have the richest environments (1542–016 (see Fig 12) & 0917+459) transpire (see Table 3) to be two out of the three RQs with \( r_e > 20 \) kpc. Moreover, both the values of \( B_{qg} \) for these objects (\( \pm 1500 \) and \( \pm 1000 \)) and their half-light radii both point towards the same conclusion (given Table 6), namely that these most massive hosts are the first- or second-ranked BCGs in clusters of richness Abell Class 0–1. A similar remarkably clean correspondence can be found within the RLQ sample. The two RLQs found by McLure & Dunlop (2000) to have the richest environments (2349–014 and 0137+012) transpire (see Table 3) to be the only two RLQs with hosts for which \( r_e > 14 \) kpc. However, no such clean correspondence is evident within the RQ sample. This may mean that RQQs, even when found in cluster environments, are less likely to be central cluster galaxies than their radio-loud counterparts. However, given the large errors on \( B_{qg} \) such a conclusion may perhaps be premature.

In summary, this comparison of the properties of the hosts of powerful AGN and bright cluster galaxies supports the basic conclusion reached by McLure & Dunlop (2000) via the completely independent measure of clustering amplitude. The larger AGN hosts have properties comparable to first or second ranked BCGs in Abell class 0→1 clusters. Both radio-loud and radio-quiet quasars thus appear to occur in a range of environments which is certainly consistent with the range of environments occupied by inactive ellipticals of comparable size and luminosity. While none of our host galaxies are comparable to the largest known BCGs, we must caution that this does not mean AGN avoid such environments (as has sometimes been inferred for FRII radio sources from the results of, for example, Prestage & Peacock 1988). In fact, because Abell Class 3 and 4 clusters are very rare, with a number density less than \( < 1.6 \times 10^{-6} \text{ h}^3 \text{Mpc}^{-3} \) (Croft et al. 1997), even an all-sky survey covering the redshift band of this study \( (0.1 < z < 0.25) \) would be expected to contain \( < 1500 \) such rich clusters. Consequently, given that only 1 in \( \approx 1000 - 10000 \) massive black holes appears to be active in the present-day universe (see Section 9 and Wisotzki, Kuhlbrodt & Jahnke 2001), we should not be surprised by the fact that no low-redshift quasar has yet been found to lie within such a rich environment or hosted by an ultra-massive BCG.

7.4 Interactions

It is often stated in the literature (eg. Smith et al. 1986, Hutchings & Neff 1992, Bahcall et al. 1997) that morphological disturbance is a common feature of the host galaxies.
of powerful AGN. In fact, at first sight, the $R$-band images presented in Appendix A and in McLure et al. (1999) show a relatively low occurrence of morphological disturbance, with only three objects ($1012+008$, $2349+014$ & $0157+001$) undergoing obvious, large-scale, interactions. However, as can be seen from the model-subtracted images, the removal of the best-fitting axisymmetric model for each host galaxy does reveal the presence of peculiarities such as excess flux, tidal tails and close companions, at lower surface-brightness levels. It is tempting to conclude from this that many of these AGN may have been triggered into action by recent interaction with companion objects. However, the true significance of the levels of interaction seen in the AGN host galaxies can only be judged by comparison with an equally-detailed investigation of the morphologies of a control sample of comparable, inactive galaxies. Here we attempt to quantify the significance of such peculiarities by constructing a control sample based on the twenty bright cluster galaxies taken from the HST archive discussed in the previous subsection.

With the range of redshifts covered by the five Abell clusters it was possible to construct 19 object pairs, each consisting of one of our AGN sample paired with a suitable control, drawn from the 20 cluster galaxies. Model-subtracted images of each of the galaxy pairs were then assessed in a blind test for the occurrence of four possible indicators of disturbance or interaction, i.e:

- Residual tidal/spiral arm features,
- Residual asymmetric flux,
- Residual symmetric flux,
- Companion object inside 10" radius,

where neither the identity of the AGN, nor which object in each pair was active, was known to the decision maker. The results of this process are shown in Table 7.

The one clear result of this test is that, taken as a group, the active galaxies display a greater occurrence of tidal/spiral arm features in their model-subtracted images than the control sample. There is no significant difference in the other three indicators. It is noteworthy however, that of the six occurrences of tidal/spiral arms, only one of these ($2349+014$) is a radio-loud object. To test whether or not this indicates an inherent difference between the radio-loud and radio-quiet AGN, the blind test was repeated for all 33 HST objects, without control galaxies, but still keeping the identity of each AGN secret, in order to prevent any subconscious bias. The results of this test are presented in Table 8.

Several features of the results presented in Table 8 merit comment. First, it can be seen that there is no significant difference between the RG and RLQ sub-samples, consistent with all the other results presented in this paper. Second, it would initially appear that there is a significant difference between the radio-loud objects and the RQQ sub-sample in terms of the occurrence of tidal/spiral arm features. However, an examination of which of the RQQs are contributing to this result shows that four of them are the objects which have combined disc/bulge model fits. Therefore, if the comparison is restricted to include only those RQQs with single-component elliptical host galaxies, this apparent difference disappears.

Table 6. The model results for the four brightest galaxies in each of five clusters obtained from the HST archive, spanning a range in richness from Abell Class 0 to Abell Class 4. The redshifts of the clusters are 0.11, 0.19, 0.23, 0.175 & 0.171 respectively. The images of A0018, A0103 and A2390 were taken through the F702W (wide $R$) filter. The images of A2390 and A1689 were taken through the F814 ($I$-band) filter. The conversion between $I$- and $R$-magnitudes has been performed assuming $R-I=0.8$ (Fukugita et al.). The final reduced images had signal-to-noise comparable to the HST host-galaxy data.

| Galaxy | $r_e$/kpc | $M_R$ | $r_e$/kpc | $M_R$ | $r_e$/kpc | $M_R$ | $r_e$/kpc | $M_R$ | $r_e$/kpc | $M_R$ |
|--------|-----------|-------|-----------|-------|-----------|-------|-----------|-------|-----------|-------|
| Abell 0018 | 30.5 | -22.65 | 24.9 | -23.62 | 27.8 | -24.52 | 75.8 | -25.33 | 100.5 | -25.51 |
| Abell 0103 | 25.4 | -22.50 | 3.4 | -22.74 | 21.2 | -23.14 | 59.0 | -24.70 | 51.9 | -24.54 |
| Abell 2390 | 16.6 | -21.47 | 7.6 | -22.69 | 12.1 | -22.99 | 18.7 | -23.62 | 35.7 | -23.68 |
| Abell 1689 | 5.2 | -21.14 | 7.6 | -22.69 | 4.1 | -22.21 | 7.6 | -23.50 | 28.8 | -23.29 |
| Abell 2218 | 22.50 | 3.4 | 22.74 | 21.2 | 23.14 | 59.0 | 23.62 | 35.7 | 23.68 |

Figure 12. The full WF2 image of the radio galaxy 1342–016 covering an area of 80" x 80". A large number of companion objects can be seen which, if at the redshift of the quasar, would appear consistent with a moderately rich cluster environment.
Table 7. The results of the blind test to determine any difference in the prevalence of features indicative of interaction or disturbance in the AGN sample (active) compared to inactive galaxies, using 20 BCG as an inactive control sample (see text for details).

| Object Type | Tidal/Spiral | Asymmetric Flux | Symmetric Flux | Companion |
|-------------|--------------|----------------|---------------|-----------|
| Active      | 5            | 8              | 5             | 7         |
| Inactive    | 0            | 10             | 7             | 9         |

Table 8. The results of the blind test to determine if any differences exist in the occurrence of features indicative of interaction or disturbance between the AGN sub-samples (see text for details).

| Object Type | Tidal/Spiral | Asymmetric Flux | Symmetric Flux | Companion |
|-------------|--------------|----------------|---------------|-----------|
| RQQ         | 6            | 5              | 1             | 2         |
| RLQ         | 1            | 6              | 2             | 4         |
| RG          | 0            | 4              | 2             | 3         |

The results of these blind comparison tests thus support two separate conclusions. First, the finding that there are no detectable differences between the elliptical AGN host galaxies and the BCG control sample provides further evidence that these two classes of object are directly comparable. Second, the apparently more frequent occurrence of tidal/spiral arm features in the RQQ sub-sample is only a manifestation of the previously-determined fact that a minority of RQQ host galaxies, two of which are not actually true quasars, have substantial disc components, and thus does not provide evidence for an inherent difference in interaction rates between radio-quiet and radio-loud objects.

Perhaps one should not be surprised by the lack of evidence for spectacular large-scale ULIRG-like interactions in our quasar sample, given that as described below, the activation of a quasar requires the delivery of only $\sim 1 \, M_\odot \, \text{yr}^{-1}$ to the central black hole.

8 THE AGN-HOST CONNECTION

As discussed by McLure et al. (1999), our finding that the hosts of quasars are almost all spheroidal galaxies allows exploitation of the now well-established correlation between black-hole mass and spheroid mass to estimate the masses of the central black holes in these objects, completely independent of any of their observed nuclear properties. We can then use this information to explore which, if any, of the observed properties of the AGN are linked to black-hole mass.

In this section we revisit this issue for a number of reasons. First, of course, our matched sub-samples are now complete. Second, in the intervening two years the application of 3-integral models has resulted in a revising down of the constant of proportionality between $M_{bh}$ and $M_{sph}$ from 0.006 (Magorrian et al. 1998) to $\sim 0.003 - 0.001$ (Kormendy & Gebhardt 2001). Third, independent estimates of the black-hole mass in a substantial number of our AGN have now been made on the basis of nuclear $H_\beta$ line-width by McLure & Dunlop (2001). Fourth, the new VLA data presented in Section 4 mean that we can now explore the relationship between black-hole mass and radio luminosity without being hampered by a large number of upper limits within the ‘radio-quiet’ subsample.

8.1 The Black-Hole Spheroid Connection

As summarized by Kormendy & Gebhardt (2001), ‘reliable’ black-hole mass estimates are now available for at least 37 nearby galaxies. As a result of this now substantial sample, clear correlations have been uncovered between black-hole mass and host-spheroid luminosity, and between black-hole mass and host stellar velocity dispersion. It has been claimed that the latter is the tighter correlation, indicating that the basic physical link is with spheroid mass. Since in the present study we do not possess measurements of host stellar velocity dispersions, we obviously have to estimate host spheroid mass on the basis of the spheroid luminosity returned by our 2-D modelling. However, a recent re-analysis of the bulge luminosities of low-redshift galaxies indicates that the relationship between bulge luminosity and black-hole mass is in fact just as tight as that between black-hole mass and central galaxy velocity dispersion (McLure & Dunlop 2002).

To estimate black-hole masses on the basis of host spheroid luminosities we have therefore adopted the following relations, while acknowledging the significant scatter present in both.

$$M_{sph} = 0.00123 L^{1.31}$$

$$M_{bh} = 0.0025 M_{sph}$$

consistent with the mass-to-light ratio relation for ellipticals determined by Jorgensen, Franx & Kjaergaard (1996) and the results of Gebhardt et al. (2000). We note that while these conversion factors are certainly also consistent with those quoted in the recent review by Kormendy & Gebhardt (2001), there is now a growing body of evidence that the true conversion factor between $M_{sph}$ and $M_{bh}$ may be a factor $\sim 2$ lower than that adopted here (Merritt & Ferrarese 2001, McLure & Dunlop 2002). If one wishes to adopt this lower factor (i.e. $M_{bh} = 0.0013 M_{sph}$) then the estimates of black-hole mass utilised in the analyses that follow in the remainder of this paper can simply be divided by a factor.
of 2. Obviously none of the arguments concerning relative masses are affected by this uncertainty, and indeed none of the physical arguments are significantly altered by such a further modest reduction in the black-hole mass estimates.

The results of applying these relations to calculate inferred bulge and black-hole masses are presented in Table 9. Also given in Table 9 (column 4) are black-hole masses derived for the majority of the quasars in these samples by McLure & Dunlop (2001) on the basis of assuming that the $H_{\beta}$-producing broad-line clouds are in Keplerian orbits.

The new spheroid-based black-hole estimates are compared with the $H_{\beta}$-derived values in Fig 13. This diagram shows that, with few exceptions, there is very good agreement ($\lesssim 0.2$ dex) between the two completely-independent estimates of black-hole mass. Indeed, given the well-known uncertainties in both approaches, the level of agreement demonstrated by Fig 13 can be viewed as adding considerable credence to the estimated values of black-hole mass, and indeed also provide support for the basic premise of a gravitationally-bound broad-line region as assumed by McLure & Dunlop (2001, 2002). The fact that four objects (three visible on the diagram) have $H_{\beta}$-based black-hole masses which lie well below the corresponding host-spheroid derived values is not really surprising since the $H_{\beta}$ line-width analysis is clearly capable of yielding a severe under-estimate of black-hole mass for any object whose broad-line region is orbiting close to the plane of the sky (but should be incapable of yielding a comparably serious over-estimate of black-hole mass). Thus for 2247+140, 0736+017, 0157+001, and 1012+008 there is a good reason for trusting the black-hole mass estimates of $\lesssim 10^9 M_\odot$ given in column 3 of Table 9, more than the lower values yielded by the $H_{\beta}$ analysis in column 4.

Accepting this, and excluding the two low-luminosity RQQs, the third column of Table 9 shows that a black hole of minimum mass $M_{bh} \geq 5 \times 10^8 M_\odot$ appears to be required to produce a nuclear luminosity corresponding to $M_R < -23$.

Beyond confirming the basic plausibility of the black-hole mass estimates, the most interesting result contained in Fig 13 is that, as suggested by the initial results of McLure et al. (1999), the black-hole masses in the radio-loud quasars are systematically larger than in their radio-quiet counterparts. In terms of summary statistics of central tendency, this difference, although clear, might not appear very dramatic – specifically, excluding the two low-luminosity RQQs:

$$\langle M_{bh} / 10^9 M_\odot \rangle = 1.67 \pm 0.36 \quad \text{median} = 1.3 \quad \text{(RG)}$$
$$\langle M_{bh} / 10^9 M_\odot \rangle = 1.64 \pm 0.22 \quad \text{median} = 1.4 \quad \text{(RLQ)}$$
$$\langle M_{bh} / 10^9 M_\odot \rangle = 1.05 \pm 0.24 \quad \text{median} = 0.8 \quad \text{(RQQ)}$$

However, the difference becomes more striking when one notices that there is an apparent threshold mass of $\approx 10^9 M_\odot$ above which lie 8/10 RGs and 10/10 RLQs, but below which lie 9/13 of the RQQs. In other words, using the black-hole mass estimation procedure outlined above, at least within this particular sample it appears that a black-hole mass $> 10^9 M_\odot$ is a necessary (albeit perhaps not sufficient) condition for the production of a powerful radio source. As measured via the KS test, this difference between the black-hole mass distributions of the radio-loud and radio-quiet subsamples is significant at the 3-$\sigma$ level, a significance which must be taken seriously given the almost perfect matching of the distributions of nuclear optical output demonstrated in Fig 4. Further support for its reality comes from the $H_{\beta}$ study by McLure & Dunlop (2001) already referred to above, in which it was found that 11/13 RLQs had $M_{bh} > 10^8.8 M_\odot$ while only 4/18 RQQs had black-hole masses above this threshold. A similar conclusion was also reached by Laor (2000) from a study of the virial masses of PG quasars.

An inevitable consequence of the matching of nuclear optical luminosities demonstrated in Fig 4, and the systematic offset in black-hole mass described above is that, on average, the black holes at the heart of the RQQs in our sample must, on average, be emitting more efficiently than their radio-quiet counterparts. This is quantified in the last two columns of Table 9 which give estimated Eddington luminosity (converted to $R$-band absolute magnitude) and then emitting efficiency $L/L_{Edd}$. The potential $R$-band Eddington luminosity of each quasar was derived in the following manner. Initially, the host-galaxy bulge luminosities provided by the two-dimensional modelling were converted into a spheroidal mass estimate using the Jørgensen et al. mass-to-light ratio (Eqn 7). The linear relation between black-hole mass and spheroidal mass (Eqn 8) was then used to estimate the central black-hole mass and resulting bolometric Eddington luminosity of each quasar. To calculate the corresponding $R$-band magnitudes it was assumed that, on average, the bolometric luminosity of a quasar can be estimated via $L_{bol} \approx 10\lambda L_{5100}$ (Laor et al. 1997), where $L_{5100}$ is the absolute monochromatic luminosity at 5100 Å. The resulting estimate of the absolute luminosity at 5100 Å was then transformed to the central wavelength of the

![Figure 13](image-url)
Table 9. The results of estimating black-hole mass from the $R$-band luminosity of the host spheroid. Columns two and three list the predicted galaxy spheroid mass and central black-hole mass respectively. For comparison with the results given in Column 3, Column 4 gives the black-hole mass estimates as derived from the $H\beta$ line width by McLure & Dunlop (2001). Column 5 lists the predicted absolute $R$-band Eddington luminosity of the black-hole, calculated according to the prescription given in Section 8.1. Column 6 gives the ratio of the predicted Eddington luminosity to the best-fitting nuclear model component.

| Source | $m_{ph}/10^{11}M_\odot$ | $m_{bh}/10^9M_\odot$ | $m_{bh-H\beta}/10^9M_\odot$ | $M_R(\text{Eddington})$ | $L_{nuc}/L_{edd}$ |
|--------|-------------------------|------------------------|-----------------------------|---------------------------|-------------------|
| RG     |                         |                        |                             |                           |                   |
| 0230−027 | 4.8                    | 1.2                    |                             |                           |                   |
| 0307+169 | 7.9                    | 2.0                    |                             |                           |                   |
| 0345+337 | 2.6                    | 0.7                    |                             |                           |                   |
| 0917+459 | 10.6                   | 2.6                    |                             |                           |                   |
| 0958+291 | 3.8                    | 0.9                    |                             |                           |                   |
| 1215+013 | 2.1                    | 0.5                    |                             |                           |                   |
| 1215−033 | 3.8                    | 0.9                    |                             |                           |                   |
| 1330+022 | 5.8                    | 1.4                    |                             |                           |                   |
| 1342−016 | 16.7                   | 4.2                    |                             |                           |                   |
| 2141+279 | 9.3                    | 2.3                    |                             |                           |                   |
| RLQ    |                         |                        |                             |                           |                   |
| 0137+012 | 8.7                    | 2.2                    | 1.1                         | −28.2                     | 0.013             |
| 0736+017 | 5.0                    | 1.3                    | 0.3                         | −27.6                     | 0.035             |
| 1004+130 | 9.4                    | 2.3                    | 3.7                         | −28.3                     | 0.089             |
| 1020−103 | 3.8                    | 1.0                    | 0.7                         | −27.4                     | 0.028             |
| 1217+023 | 5.9                    | 1.5                    | 0.8                         | −27.8                     | 0.043             |
| 2135−147 | 3.9                    | 1.0                    | 2.6                         | −27.4                     | 0.049             |
| 2141+175 | 4.2                    | 1.1                    | 1.7                         | −27.5                     | 0.061             |
| 2247+140 | 6.5                    | 1.6                    | 0.1                         | −27.9                     | 0.022             |
| 2349−014 | 12.2                   | 3.1                    | 1.8                         | −28.6                     | 0.014             |
| 2355−082 | 5.3                    | 1.3                    | 0.7                         | −27.7                     | 0.014             |
| RQQ    |                         |                        |                             |                           |                   |
| 0052+251 | 2.3                    | 0.6                    | 0.6                         | −26.8                     | 0.107             |
| 0054+144 | 5.2                    | 1.3                    | 2.5                         | −27.7                     | 0.052             |
| 0157+001 | 12.2                   | 3.1                    | 0.2                         | −28.6                     | 0.011             |
| 0204+292 | 3.6                    | 0.9                    | 0.6                         | −27.3                     | 0.020             |
| 0244+194 | 1.9                    | 0.5                    | 0.3                         | −26.6                     | 0.046             |
| 0257+024 | 0.4                    | 0.1                    |                             | −24.9                     | 0.008             |
| 0923+201 | 3.4                    | 0.8                    | 2.6                         | −27.2                     | 0.088             |
| 0953+414 | 2.1                    | 0.5                    | 0.7                         | −26.6                     | 0.355             |
| 1012+008 | 6.4                    | 1.6                    | 0.02                        | −27.9                     | 0.026             |
| 1549+203 | 1.9                    | 0.3                    | −25.9                       | 0.169                      |
| 1635+119 | 3.2                    | 0.7                    | 0.4                         | −26.9                     | 0.007             |
| 2215−037 | 4.7                    | 1.2                    | −27.6                       | 0.010                      |
| 2344+184 | 1.3                    | 0.3                    | −26.8                       | 0.002                      |

$R-$ filter (6500 Å) assuming a power-law quasar spectrum ($\alpha = 0.2$). Using this prescription, a quasar with a central black-hole mass of $10^9 M_\odot$ will have an Eddington-limit absolute $R-$band magnitude of $M_R(\text{Edd}) = −27.5$.

8.2 The Black-Hole–Radio Connection

When the matched samples of RLQs and RQQs were first defined, all that was known about the radio properties of the RQQs was that their $P_{5\text{GHz}} \approx 10^{24}$ WHz$^{-1}$sr$^{-1}$ (Dunlop et al. 1993). However, with the completion of the VLA detection programme described in Section 4 we can now explore how the radio properties of all the AGN in our combined sample relate to their estimated black-hole masses.

In Fig 15 total radio luminosity, $P_{5\text{GHz}}$, is plotted against estimated black-hole mass for all the objects in our combined sample (with the single exception of 1549+203, which was not observed with the VLA to the same depth as the other objects). Also included in this diagram are the appropriate points for 8 nearby galaxies taken from Franceschini et al. (1998), with the Milky Way and M87 highlighted.
Several significant trends are evident in this diagram. First it can be seen that, within the RQQ sample, radio luminosity correlates, albeit weakly, with black-hole mass, although our sample doesn’t span a wide dynamic range and the correlation is not formally significant. Second, and perhaps more interestingly, the RQQs are consistent with the locus described by the nearby galaxies (for which, of course, the black-hole masses are based directly on stellar dynamics). This simultaneously reinforces the plausibility of our black-hole mass estimates, and suggests a common physical origin for the radio emission from RQQs and nearby optically more-quiescent bulges. Whatever this physical mechanism is, it must explain the observed strong dependence of minimum radio output on black-hole mass, in which the power-law index $\gamma$ is at least 2 (adopting $P_{\text{radio}} \propto M_{\text{bh}}^\gamma$). As noted by Franceschini et al. (1998), $\gamma = 2.2$ is the prediction of ADAF models, while $\gamma \approx 2$ is a generic prediction of any mechanism in which radio output is proportional to the area of the accretion disc.

Our results therefore provide further support for previous claims that there is a minimum radio output from black holes and that this is a strong function of mass. However, while this obviously implies that the most massive black holes can never be very radio-quiet, Fig 15 graphically demonstrates that the genuine radio-loud objects lie on a quite distinct relation. Interestingly, our radio-loud sample also displays an (albeit weak) internal correlation between $P_{\text{GHI}}$ and estimated black-hole mass. Either this relationship must be much steeper ($P \propto M_{\text{bh}}^5$) than the minimum radio-output relation, or, as illustrated in Fig 15, the black-hole mass dependence of the upper envelope may be consistent with that of the lower envelope but offset by several orders of magnitude.

In fact the findings of Lacy et al. (2001) strongly favour the latter option. They report a relationship between the logarithm of radio power and black-hole mass which has a slope of $1.4 \pm 0.2$. However, inspection of their Fig 2 shows that their data are also consistent with an upper and lower envelope with slope $\approx 2.5$ as illustrated here in Fig 15.

We note here that it has recently been claimed by Oshlack, Webster & Whiting (2001) that the radio-loud Seyfert galaxy PKS 2004–447 contradicts the existence of such an envelope, and shows that relatively low-mass black holes can produce a powerful radio source. However, this claim is based on their black-hole mass estimate for this object of $\approx 5 \times 10^6 M_\odot$ and it seems almost certain that this value, derived from the relatively narrow emission lines displayed by PKS 2004–447, is a serious underestimate of the true mass. The reason for this is that since this object displays optically variability and has a compact radio source, it seems very likely that its orbiting broad-line region lies close to plane of the sky. In this situation (as found in our own sample for, e.g. 1012+008 – see Table 9) the black-hole mass derived from $H_\beta$ emission-line width can under-estimate the true value by a factor $> 100$. In Fig 15 we have plotted the position of PKS 2004–447 after re-calculating its black-hole mass on the assumption that the orbital plane of its broad-line region is oriented within $\approx 5$ degrees of the plane of the sky (see McLure & Dunlop 2001). This increases the estimated black-hole mass to $\approx 4 \times 10^6 M_\odot$, and moves PKS 2004–447 into complete consistency with the upper envelope delineated by the other available data. We have not corrected the observed radio emission for doppler boosting - if we did then the this object would simply move downwards in Fig 15.

Thus both our own results and other currently available data support the existence of both a lower and an upper envelope to the radio luminosity which can be produced by a black hole of given mass. Indeed Fig 15 serves to provide an elegant explanation not only of our own findings, but also of several other well-known results in the literature.

First, because we selected our radio-loud objects on the basis that their radio luminosities were greater than $log_{10}(P/\text{WHz}^{-1}\text{sr}^{-1}) = 24$, the form of the upper envelope in these figures can be seen to be perfectly consistent with our finding that virtually all our radio-loud objects have $M_{\text{bh}} > 10^7 M_\odot$. In the context of this diagram this finding can be seen to be a result of our radio-selection criterion rather than the existence of some critical black-hole mass for the production of relativistic jets within a galaxy mass halo.

Second, the steep dependence of this apparent upper envelope on black hole mass means that radio sources up to $log_{10}(P/\text{WHz}^{-1}\text{sr}^{-1}) = 27 - 28$, consistent with the most luminous 3CR galaxies even at $z = 1$ can be produced by black holes (and hence hosts) only a factor of 2-3 more massive, and most likely will be due to the steep high-mass slope of the black-hole mass function. This provides a natural explanation of why radio galaxies spanning the top two orders
of magnitude in radio power all appear to be giant ellipticals with stellar masses which agree to within a factor of $\lesssim 2$ (Jarvis et al. 2001).

Third, from Fig 15 it can be seen that the properties of AGN hosts will inevitably become more mixed if the radio-luminosity threshold is move down to, say, $P_{5GHz} \lesssim 10^{22}$ W Hz$^{-1}$sr$^{-1}$. Given the lower boundary in Fig 15, radio sources in this power range can still be produced by black holes as massive as $10^9 M_\odot$ living in giant ellipticals, but it is also clear that black holes with masses as low as $10^{7.5} M_\odot$ are capable of producing detectable radio sources at this level. As demonstrated by the properties of the lower-luminosity objects in our own sample, and by the studies of ‘normal’ galaxies at low redshift, such black holes can be housed in disc-dominated galaxies as well as discless lower-mass spheroids.

8.3 The origin of radio loudness

While the broad dependence of radio luminosity on black-hole mass may explain the relative numbers of radio-quiet and radio-loud objects, it is also clear from Fig 15 that black hole mass cannot be invoked to explain the range of radio luminosities displayed by quasars of very similar optical luminosity. Specifically it is clear from this figure that objects powered by equally massive black holes can differ in radio luminosity by $\lesssim 4$ orders of magnitude.

Interestingly Fig 15 also allows us to rule out the possibility that the radio output of a black hole of given mass is driven simply by fuel supply. This is because, while it has long been known that RQQs are not radio silent, it can be seen from Fig 15 that several of the RQQs in our sample are in fact as radio-silent as they could possibly be, given their estimated black-hole masses. In other words at least some of the black holes in RQQs produce as little radio output as their counterparts in completely quiescent nearby galaxies, despite the fact that black holes in quasars are clearly in receipt of sufficient fuel to produce powerful optical emission.

Thus, while the results of this study cannot by themselves provide a definitive answer to the long standing question of the origin of radio loudness, they can be used to focus the argument by excluding several possible explanations. In summary, our results can be used to exclude host galaxy morphology, black-hole mass, or black hole fueling rate as the primary physical cause of radio loudness.

We are thus forced to the conclusion that the production of powerful relativistic radio jets is driven by some other property of the black hole itself, the most likely candidate being black-hole spin (e.g. Blandford 2000, Wilson & Colbert 1995). Indeed, following the suggestion by Blandford (2000), we speculate that the evolutionary track followed by an activated black hole may manifest itself as a near vertical descent in Fig 15, with an active rapidly spinning hole appearing first as a powerful radio source close to the upper envelope in this diagram, and then descending towards the lower envelope as the spin energy of the hole is exhausted.

9 COSMOLOGICAL IMPLICATIONS

9.1 Relative numbers of radio-loud and radio-quiet quasars

One interesting aspect of the different black-hole mass thresholds for RQQs and RLQs uncovered by this study is that this difference of a factor of 2 in minimum black-hole mass provides a natural explanation of why radio-quiet quasars outnumber their radio-loud counterparts by a factor of $\lesssim 10$ (and why this factor may reduce with increasing optical luminosity; Hooper et al. 1995, Goldschmidt et al. 1999).

This is because if our black-hole mass thresholds of $M_{bh} > 5 \times 10^8 M_\odot$ for RQQs, and $M_{bh} > 1 \times 10^9 M_\odot$ for radio-loud objects are converted back into absolute spheroid
magnitudes in the K-band (using the absolute R-band magnitudes given in Table 4, and then converting into K adopting \( R - K = 2.5 \) as justified by the results in Section 6.6) we find that RQQs can be hosted by spheroids with \( M_K < -25.3 \) while radio-loud objects require a host spheroid with \( M_K < -26 \). With reference to the K-band luminosity function these absolute magnitude limits correspond to \( L > 1.5L^* \) and \( L > 3L^* \) respectively (Gardner et al. 1997, Szokoly et al. 1998, Kochanek et al. 2001) which leads to the prediction that, in the present day universe there are \( 1 \times 10^{-4} \) galaxies per Mpc\(^3\) capable of hosting an RQQ brighter than \( M_V = -23 \), but only \( 1 \times 10^{-5} \) galaxies per Mpc\(^3\) are capable of hosting an RLQ. This means that, assuming low-redshift quasars arise from a randomly-triggered subset of massive black-holes in spheroids, the factor of two difference in black-hole mass threshold translates into a factor of 10 difference in the number density of RQQs and RLQs due simply to the steepness of the bright end of the spheroid luminosity function.

9.2 Black-hole activation fraction at \( z \approx 0 \)

We can go further and use the above numbers to estimate the activation fraction for massive black holes in the local universe. First we note that the radio-loud objects in our sample have \( P_{22\,\text{GHz}} > 1 \times 10^{25} \text{ WHz}^{-1}\text{sr}^{-1} \) (assuming a spectral index \( \alpha \approx 0.8 \)). Reference to the local radio luminosity function in Dunlop & Peacock (1990), or to the more recent determination at 1.4 GHz from the 2dF by Sadler et al. (2001), indicates that the present-day number density of radio sources above this threshold is \( \approx 1 \times 10^{-8} \text{ Mpc}^{-3} \). Thus, in the present-day universe we find that 1 in every 1000 black holes more massive than \( 1 \times 10^9 \text{ M}_\odot \) (or equivalently spheroids more luminous than \( 3L^* \)) is radio active in the present-day universe.

Turning to the RQQ population, the present-day number density of quasars with \( M_V < -23 \) can be estimated by extrapolating the QSO OLF deduced at \( z \approx 0.4 \) by Boyle et al. (2000) back to \( z = 0 \) assuming luminosity evolution \( \propto (1+z)^3 \). This produces an estimated number density of \( \approx 5 \times 10^{-8} \text{ Mpc}^{-3} \), which is at least consistent with the direct determination of the low-z OLF attempted by Londish, Boyle & Schade (2000). Boosting this number by a factor of two to allow for obscured quasars (i.e. adopting an opening angle of \( \approx 45^\circ \) for consistency with the correction factor which applies to active nuclei) then leads to the conclusion that 1 in every 1000 black holes more massive than \( 5 \times 10^8 \text{ M}_\odot \) (or equivalently spheroids more luminous than \( 1.5L^* \)) is producing quasar-level optical activity.

Thus both these comparisons yield the same result, namely that in the present-day (\( z < 0.1 \)) universe 1 in every 1000 massive black holes is active, and that the reason RQQs outnumber RLQs by a factor of 10 is a direct result of the fact that their respective minimum black-hole masses differ by a factor of two.

We note here that Wisotzki, Kuhlbrodt & Jahnke (2001) have recently concluded that the present-day black-hole activation fraction is 1 in 10000, or 1 in 5000 if the quasar LF is boosted by a factor of two as assumed above. However, reference to their Fig 2 shows that this factor results from adopting the spheroid luminosity function of Lin et al. (1999). If instead Wisotzki et al. were to adopt the Kochanek et al. (2001) luminosity function then it is clear that (at least for the mass range probed by the present study) Wisotzki et al.’s analysis would also yield an activation fraction of 0.1%.

9.3 Black-hole activation fraction at \( z \approx 2 - 3 \)

At the peak epoch of quasar activity between \( z \approx 2 \) and \( z \approx 3 \), the co-moving number density of powerful radio sources with \( P_{22\,\text{GHz}} > 1 \times 10^{25} \text{ WHz}^{-1}\text{sr}^{-1} \) has risen to \( \approx 1 \times 10^{-6} \text{ Mpc}^{-3} \) (Dunlop & Peacock 1990) while that of optically-selected quasars brighter than \( M_V = -23 \) is \( \approx 5 \times 10^{-6} \text{ Mpc}^{-3} \) (Warren, Hewett & Osmer 1995, Boyle et al. 2000).Doubling the latter figure to make the same correction for obscured RQQs as applied in the previous subsection at low redshift, leads to the conclusion that, as at \( z \approx 0 \), RQQs at \( z \approx 2.5 \) outnumber their radio-loud counterparts by a factor of \( \approx 10 \). However, at \( z \approx 2.5 \) the co-moving number density of both classes of object is enhanced by a factor of 100, giving an activation fraction of massive black holes of \( \approx 10\% \), assuming that the entire massive black-hole population is in place by that epoch.

9.4 The cosmological evolution of AGN

A renewed attempt to model the cosmological evolution of AGN in the light of our host-galaxy results is obviously beyond the scope of this paper. However, if one assumes that the same mass thresholds uncovered in our low-redshift study apply at high redshift, then one simple interpretation of the dramatic (and very similar) evolution of radio-loud and radio-quiet AGN between \( z \approx 2.5 \) and the present-day is that it is primarily due to density evolution, with the active fraction of massive black holes dropping from 10% to 0.1% over this period. There must of course also be an element of luminosity evolution because pure density evolution was excluded long ago as an acceptable representation of the evolving RLF or OLF. However, as suggested by Miller, Percival & Lambert (in preparation) if individual quasars produce declining light curves it is inevitable that a larger fraction will be observed closer to peak luminosity as the activation rate rises, and Fig 14 indicates that at least some of the host galaxies uncovered in the present study should be capable of producing quasars with \( M_V \approx -28 \) if observed while accreting at the Eddington limit.

One prediction of this scenario is that, as a larger fraction of massive black holes become more efficiently fueled, the apparent mass difference between RLQ and RQQ hosts should grow with increasing redshift. The reason for this is that, if fueling rates are, on average, somewhat higher at high redshift we can expect optically-selected QSOs of a given luminosity to be produced by black holes with, on average, lower masses than at low redshift. However, if the \( m_{\text{th}} > 10^9 \text{ M}_\odot \) mass threshold for powerful radio activity found in this study continues to apply at high redshift, then radio selection will continue to yield only the most massive black holes residing (presumably) in the most massive spheroids.

In fact, studies of the hosts of RLQs and RQQs out to \( z \approx 2 \) are already uncovering evidence of just such a trend.
Specifically, Kukula et al. (2001) have used NICMOS on the HST to measure the rest-frame $R$-band luminosities of the hosts of matched samples of RLQs and RQQs at $z \simeq 1$ and $z \simeq 2$. Kukula et al. find that the hosts of RLQs are essentially unchanged in mass over this redshift range, suggesting that the minimum mass threshold for powerful radio emission indicated by the upper locus in Fig 15 applies at all redshifts and is therefore of physical significance. However, they also find that the average ratio of RLQ:RQQ host luminosity rises from the value of 1.5 found here at $z \simeq 0.2$, to $\simeq 2$ at $z \simeq 1$, and $\simeq 3$ at $z \simeq 2$, suggesting a progressive drop in the average mass of RQQ hosts with increasing redshift, a finding supported by the results of other recent studies (Ridgway et al. 2000, Rix et al. 2000).

10 CONCLUSIONS

In this paper we have reported the extensive results which follow from completion of our HST imaging programme of the host galaxies of low-redshift radio-quiet quasars, radio-loud quasars, and radio galaxies. This paper represents the completion of a programme which has spanned most of the last decade, commencing with the deep infrared imaging of the same sample of objects by Dunlop et al. (1993). The depth and quality of our HST data have now allowed us to determine accurately all the basic structural parameters of the hosts of these 3 classes of powerful AGN. Because of this, and because of the wealth of data now available for this sample at other wavelengths (including the new deep VLA observations also reported here), we have been able to address a wide range of issues of which are hopefully of interest to workers in several different areas of extra-galactic research. Consequently we conclude with a detailed summary of the main conclusions of this study, structured to assist the interested reader in moving directly to the sections which may be of most relevance to their own work.

10.1 Results from analysis of the HST images

From the detailed 2-dimensional modelling of the new HST images presented here in combination with the data reported by McLure et al. (1999) we find that (as detailed in Sections 5 and 6):

- The hosts of all three classes of powerful AGN display a surface-brightness scalelength (Kormendy) relation identical (in both slope and normalization) to that displayed by inactive massive ellipticals predominantly found in clusters.
- The hosts of all three classes of powerful AGN display a distribution of axial ratio which is indistinguishable to that which has been long established for normal elliptical galaxies.

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10.2 Results incorporating infrared images

From a joint analysis of the HST optical images with existing $K$-band images of the same sample (Dunlop et al. 1993, Taylor et al. 1996, McLure, Dunlop & Kukula 2000) we find that (as detailed in subsections 6.6 and 6.7):

- The hosts of all three classes of powerful AGN are red galaxies, with typical rest-frame optical-infrared colours $R-K \simeq 2.5$. Consistent with the results of Nolan et al. (2001), these colours are as expected from an evolved stellar population of age $10-13$ Gyr, indicating that, as for quiescent massive ellipticals, the stellar mass of an AGN host galaxy is dominated by a well-evolved stellar population formed at high redshift ($z > 2$).
- The optical-infrared colours of the RQQ and RLQ nuclei are statistically indistinguishable, and lie in the range $R-K \simeq 2-4$. For the RG nuclei, the relation between $R-K$ colour and $R$-band luminosity is as expected under the assumption that they all contain obscured nuclei with intrinsic optical luminosities comparable to the quasars, reddened by dust with $\kappa_\lambda \propto \lambda^{0.95}$. This result strongly favours unification of RGs and RLQs via orientation rather than time.

10.3 Relation of quasar hosts to normal galaxies

From a comparison of our results with those derived from recent studies of the ‘normal’ massive galaxy population, and of low-redshift ULIRGs and Seyferts, we find that (as detailed in Section 7):

- The bulge:disc ratio of AGN hosts is a function of AGN luminosity, with disc-dominated hosts dying out above the traditional quasar:Seyfert boundary at $M_V(nuc) \simeq -23$.
- In contrast to ULIRGs, quasar hosts lie in a region of the fundamental plane (as judged from the photometric projection of the FP offered by the Kormendy relation) which has been shown to be occupied by the most massive and apparently old population of ellipticals which display ‘boxy’ isophotes and distinct kinematic cores. This provides a strong argument against the possibility that the $r^{1/4}$-law luminosity profiles displayed by the vast majority of the quasar hosts in our sample are the result of recent mergers between massive disc galaxies. This result also offers little support for a strong evolutionary connection between ULIRGS and quasars in the low-redshift universe.
- The basic structural properties of AGN hosts are indistinguishable from those of inactive brightest cluster galaxies. Consistent with the environmental study of McLure & Dunlop (2000), the largest host galaxies in our sample are of comparable mass to the BCGs found at the heart of Abell Class 1 or Class 2 clusters. However, the lack of any host.
galaxies as massive as an Abell Class 4 cluster BCG is as expected given the relative rarity of such rich environments. All the evidence considered is consistent with the AGN hosts being drawn, essentially at random, from the present-day massive-elliptical galaxy population.

- There is no statistically-significant evidence that AGN hosts display more signs of large (kpc) scale disturbance, or multiple nuclei than normal comparably-massive quiescent ellipticals.

10.4 The black-hole spheroid connection

Combining our HST data with our new VLA data, and utilising recent results on the black-hole:spheroid mass relation (Section 8),

- Estimates of central black-hole mass for the quasars in our sample based on host-galaxy spheroid luminosity agree well with those derived from Hβ emission-line width by McLure & Dunlop (2001). For three objects the latter technique yields a much lower value, but this is as expected for a small subset of objects if the broad-line region has a predominantly disc-like structure.

- Based on the relation \( M_{bh} = 0.0025 M_{sph} \), we find that all the quasars in our sample are powered by a black-hole of mass \( m_{bh} > 5 \times 10^5 M_{\odot} \), but that the radio-loud objects lie above an even higher mass threshold, with \( m_{bh} > 10^8 M_{\odot} \).

- The most efficient RQQ in our sample appears to be radiating at \( \lesssim 35\% \) of the Eddington limit, but the vast majority of these low-redshift AGN are emitting at \( \lesssim 1 \% \) of the Eddington limit. Based on these calculations, the most massive objects in this low-redshift sample appear capable (if fuelled at maximum efficiency) of producing quasars with \( M_V \approx -28 \), comparable to the most luminous objects known at high redshift.

- The black-hole mass difference between RLQs and RQQs reflects the existence an apparent upper and lower limit to the radio output that can be produced by a black hole of a given mass. Both the upper and lower thresholds on radio luminosity appear to be a strong function of black mass (\( \propto m_{bh}^{-2.5} \)). At least some RQQs in our sample appear to be as radio-weak as ‘normal’ inactive comparably-massive spheroids found in the low-redshift universe, despite the fact that must be receiving sufficient fuel to power the optical quasar nucleus.

- While there is a broad and clear trend for increasing radio luminosity with increasing black-hole mass, it is now clear that host morphology, black-hole mass, and fuel supply can each be excluded as the primary physical explanation of why a subset of quasars are up to 4 orders of magnitude more radio luminous than their ‘radio-quiet’ counterparts. Consequently we argue that black-hole spin is the most plausible (perhaps only) remaining feasible explanation for the production of a powerful radio source.

10.5 Cosmological implications

Considering the implications of our results for our understanding of the nature and evolution of AGN populations as a function of redshift, we find that (as detailed in Section 9):

- The relative numbers of radio-loud and radio-quiet quasars can be naturally explained by the above-mentioned mass thresholds, combined with the form of the bright end of the elliptical galaxy luminosity function.

- The activation fraction of massive black holes is \( \approx 0.1\% \) in the present-day universe, rising to \( \approx 10\% \) at \( z \approx 2 - 3 \), corresponding to the peak epoch of quasar activity in the universe.

- The black-hole mass threshold for powerful radio activity can explain the trend for a growing gap with increasing redshift between the masses of RLQ hosts and RQQ hosts uncovered by recent HST-based studies of high redshift quasars.

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REFERENCES

Abraham R.G., Crawford C.S., McHardy I.M., 1992, ApJ, 401, 474
Bahcall J.N., Kirhakos S., Schneider D.P., 1994, ApJ, 435, L11
Barnes J.E., Hernquist L., 1992, ARA&A, 30, 705
Barthel P.D., 1989, ApJ, 336, 506
Blandford R.D., 2000, PTRSA, in press (astro-ph/0001499)
Boyle B.J., Shanks T., Croom S.M., Smith R.J., Miller L., Loaring N., Heymans C., 2000, MNRAS, 317, 1014
Brinkmann W., Yuan W., Siebert J., 1997, A&A, 319, 413
Canalizo G., Stockton A., 1997, ApJ, 480, L5
Capaccioli M., Caon N., D’Onofrio M., 1992, MNRAS, 259, 323
Crawford C.S., Fabian A.C., 1995, MNRAS, 273, 827
Croft R.A.C., Dalton G.B., Efstathiou G., Sutherland W.J., Madgwick J.S., 1997, MNRAS, 291, 305
De Koff S., et al., 1996, ApJS, 107, 621
de Vries W.H., O’Dea C.P., Barthel P.D., Fanti C., Fanti R., Lehnert M.D., 2000, ApJ, 120, 2300
Disney M.J., et al. 1995, Nature, 376, 150
Dunlop J.S., Peacock J.A., 1990, MNRAS, 247, 19
Dunlop J.S., Taylor G.L., Hughes D.H., Robson E.I., 1993, MNRAS, 264, 455
Faber S.M., et al., 1997, AJ, 114, 1771
Fabian A.C., 1999, MNRAS, 308, L39
Franceschini A., Vercellone S., Fabian A.C., 1998, MNRAS, 297, 817
Franceschini A., Hasinger G., Miyaji T., Malquori D., 1999, MNRAS, 310, L5
Fukugita M., Shimasaku K., Ichikawa T., 1995, PASP, 107, 945
Gardner J.P., Sharples R.M., Frenk C.S., Carrasco B.E., 1997, ApJL, 480, L99

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APPENDIX A: NEW HOST-GALAXY IMAGES AND MODEL FITS

The images, two-dimensional model fits, and model subtracted images for the 14 AGN for which the new HST data are reported in this paper are presented in this appendix in Figs A1 to A14. A grey-scale/contour image of the final reduced F675W $R$-band image of each AGN is shown in the top-left panel (panel A) of each figure, covering a region of $12.5 \times 12.5$ arcsec centered on the target source. The surface-brightness level of the lowest contour is indicated in the top-right corner of this panel, and the grey-scale has been chosen to highlight structure close to this limit. Higher surface-brightness contours are spaced at intervals of $0.5$ mag. arcsec$^{-2}$, and have been super-imposed to emphasize brighter structure in the centre of the galaxy/quasar. Panel B in each figure shows the best-fitting two-dimensional model, complete with the unresolved nuclear component (after convolution with the empirical PSF described above), contoured in a manner identical to panel A. Panel C shows the best-fitting host galaxy as it would appear if the nuclear component were absent, while panel D is the residual image which results from subtraction of the full two-dimensional model (in panel B) from the raw $R$-band image (in panel A), in order to highlight the presence of any morphological peculiarities such as tidal tails, interacting companion galaxies, or secondary nuclei. Within each figure, all four panels are displayed using the same greyscale.

A1 Notes on individual objects

Here we provide a brief discussion of each of the 14 new HST images, with reference to other recent HST and ground-based data. Comparably-detailed descriptions of the other 19 objects in the sample can be found in McLure et al. (1999) and are not repeated here. Additional details on each object, along with a description of the main features of our $K$-band images, can be found in Dunlop et al. (1993) and Taylor et al. (1996). The results of off-nuclear optical spectroscopy and resulting spectral model-fitting for several of the objects in this sample can be found in Hughes et al. (2000) and Nolan et al. (2001).

A1.1 The Radio Galaxies

0230$-$027 (PKS 0230$-$027, OD$-$050)
The new $R$-band HST image shows this galaxy to be uniform and round with no sign of obvious distortion. The detail of the WF2 image shown in Fig A1 reveals $\approx 10$ apparent companion objects lying in a roughly circular formation around the galaxy, at a radius of $\approx 30 \rightarrow 60$ kpc. The brightest of these companions lies at a projected distance of $\approx 60$ kpc to the NW. The full WF2 chip reveals there to be $\approx 20$ faint companion objects within a radius of $\approx 200$ kpc.

0307$+$169 (3C 079, 4C$+$02.36)
The $R$-band image of this source shown in Fig A2 reveals it to be a classic brightest-cluster galaxy (BCG). Two bright companion objects can be seen to the North and South with three or four accompanying tidal arm features emanating from the central galaxy. The suggestion that this source is in the process of undergoing merger activity is strengthened by the overlying contours which show three distinct cores. The image of the full WF2 chip reveals several more bright companions within a radius of $\leq 100$ kpc.

This source has recently been imaged by the HST PC during the 3CR snapshot survey (De Koff et al. 1996) through the F702W (wide $R$) filter. The 280-second integration presented by De Koff et al. confirms the complex multiple structure of this object with the authors noting that the optical and radio axes are aligned to within 15 deg. McCarthy et al. (1995) imaged this source both in the $R$-band and through an $H\alpha$ emission-line filter, detecting a curving filament of extended $H\alpha$ emission stretching some 12$''$ to the NW.

1215$-$033 (PKS 1215$-$033)
The $R$-band image of this source shows it to be a uniform round galaxy with no obvious signs of interaction or disturbance. There are two companion objects lying just on the edge of the detail shown in Fig A3, to the SW and ESE at projected distances of $\approx 60$ kpc.

1215$+$013 (PKS 1215$+$013)
The $R$-band image of this source shown in Fig A4 reveals numerous faint companion objects around the central galaxy, the brightest lying $\approx 30$ kpc to the South. The overlying contours for this source suggest some sort of disturbance immediately to the North of the galaxy core.

1330$+$022 (3C 287.1, 4C$+$02.36)
Numerous companion objects are seen in the new $R$-band image of this source shown in Fig A5, the brightest of which can be seen some 25 kpc to the NE. There is a linear tidal feature $\approx 35$ kpc to the SW. The overlying contours for this source show there to be an apparent second nucleus at less than 1$''$ separation to the WNW.

This object was in the sample imaged in the $R$-band during the 3CR snapshot survey (De Koff et al. 1996). The 280-second image presented by De Koff et al. confirms the presence of the second nucleus. 3C 287.1 has been shown to have a power-law X-ray spectrum by Crawford & Fabian (1995).

1342$-$016 (PKS 1342$-$016, MRC 1342$-$016)
The new HST image shows this galaxy to be large, luminous and uniform. A bright foreground star is present $\approx 13''$ to the SWW, on the edge of the frame shown in Fig A6. No obvious companion objects are present in the detail shown in Fig A6 although the full WF2 chip image shown in Fig 13 reveals there to be a large number of companion objects, consistent with a moderately rich cluster.

A1.2 The Radio-Loud Quasars

1020$-$013 (PKS 1020$-$013, UT 1020$-$013)
The $R$-band image of this source shown in Fig A7 reveals it to have a comparatively small and faint host, with an obvious PA of $\approx 110$ degrees. There is a triangle of three faint companions directly to the South. This object was identified as an X-ray source in the ROSAT All-Sky Survey (Brinkmann et al. 1997).

1217$+$023 (PKS 1217$+$02, ON 029)
The $R$-band image of this relatively nuclear-dominated ob-
ject shown in Fig A8 clearly reveals an elliptical-looking host galaxy with an apparent PA of \( \approx 100 \text{ deg} \). There is a group of four faint apparent companion objects running North-South at a projected distance of \( \approx 40 \text{ kpc} \).

**2135−147** (PKS 2135−14, PHL 1657)

The image of this source shown in Fig A9 shows a large companion galaxy at a projected distance of \( \approx 30 \text{ kpc} \) ESE and an apparent close-in companion, or secondary nucleus, at a distance of \( \approx 10 \text{ kpc} \). A recent analysis of the spectrum of the close-in companion by Canalizo & Stockton (1997) has shown that this object is actually a foreground star. This object has also been imaged in the V-band with HST by BKS. Their analysis also reveals the underlying host to be best matched by an elliptical galaxy model. This quasar was identified as an X-ray source in the ROSAT All-Sky Survey (Brinkmann et al. 1997).

**2355−082** (PKS 2355−082, PHL 6113)

The new R-band HST image of this source in Fig A10 reveals the presence of an apparently early-type host with a PA of \( \approx 180 \text{ deg} \). A group of five small companion objects can be seen to the NE. This object was identified as an X-ray source in the ROSAT All-Sky Survey (Brinkmann et al. 1997).

### A1.3 The Radio-Quiet Quasars

**0052+251** (PG 0052+251)

Although not very obvious in the grey-scale representation used in Fig A11, the host galaxy of this quasar can be seen to have clear spiral structure in the raw R-band image. There are two spiral arms present to the East and West of the nucleus with the Eastern arm being more extended and apparently terminating in a companion object. The overlying contours for this image reveal it to have a highly-luminous nuclear component.

This object has recently been imaged in the H-band by McLeod & Rieke (1994) and in J-, H- and K-band by Hutchings & Neff (1997). Hutchings & Neff performed an analysis of the numbers, magnitudes and colours of the companion objects of this source, covering an angular extent comparable to the full WF2 image. They identify some 22 companion objects brighter than \( K = 19.6 \), the vast majority of which have colours consistent with mature stellar populations. Comparing their near-infrared images with previous optical studies, Hutchings & Neff note that there is no evidence for tidal structure made up of old stars, but conclude that the host is in a phase of secondary star formation possibly induced by interaction with one of the close group of companions. 0052+251 was also included in the V-band HST imaging programme of Bahcall, Kirhakos & Schneider, who list the host morphology as spiral and identify many of the knots seen in the Eastern arm with H I regions. Interestingly Bahcall, Kirhakos & Schneider comment that the inner regions of the surface-brightness profile of 0052+251 are well matched by an \( r^{-1.4} \) law, consistent with the best-fit elliptical host from the K-band imaging of Taylor et al. (1996), and suggestive that there may be a strong bulge component to the host galaxy. This is of course precisely what we have found through the construction of a combined disc+bulge model for the galaxy as detailed in Section 5. Despite the fact that the eye is drawn to the high surface-brightness spiral arms in the HST image, we find that in fact 70% of its total R-band light is contributed by the spheroidal component. Finally, we note that this quasar has also been identified as an X-ray source in the ROSAT All-Sky Survey (Yuan et al. 1998).

**0204+292**

In the new R-band image of this quasar the host galaxy appears to be elliptical with a well-defined position angle of \( \approx 90 \text{ deg} \). There are two companion objects on the sub-image shown in Fig A12, \( \approx 70 \text{ kpc} \) to the WSW and \( \approx 45 \text{ kpc} \) to the NW respectively. This quasar was identified as an X-ray source in the ROSAT All-Sky Survey (Yuan et al. 1998).

**1549+203** (LB 906, 1E 15498+203)

It is immediately obvious from the new R-band image of this object that the host galaxy is small, and relatively faint compared to the nuclear component. Using a different grey-scale to that used in Fig A13, there is a suggestion of spiral-like features to the NW and SE of the nucleus, with the SE arm terminating at the apparent companion object which can be seen \( \approx 20 \text{ kpc} \) to the E of the nucleus. A large, luminous elliptical galaxy can just be seen on the edge of the frame to the SW, with numerous fainter companions visible inside a radius of \( \approx 60 \text{ kpc} \). The full WF2 image would appear to show numerous companion objects although, as pointed out by Taylor et al. 1996, the density of the environment of this quasar is uncertain due to the presence of a nearby foreground cluster at \( z \approx 0.14 \) (Stocke et al. 1983). The first electronic images of this quasar were presented by Hutchings & Neff (1992). They imaged the object in both the V- and I-bands, detecting what looked like a bar structure running North-South through the nucleus, while noting that the surface-brightness profile of 1549+203 was not exponential. This quasar was identified as an X-ray source in the ROSAT All-Sky Survey (Yuan et al. 1998).

**2215−037** (EX 2215−037)

The image of this quasar shown in Fig A14 reveals the host to be compact and apparently undisturbed. A small apparently unresolved companion is detected only \( \approx 10 \text{ kpc} \) to the SW of the quasar nucleus. A large elliptical galaxy can be seen some 65 kpc to the North with a second apparently unresolved companion \( \approx 50 \text{ kpc} \) to the East. This object has been previously imaged on the HST PC using the F702W (wide R) filter by Disney et al. (1995). Using a two-dimensional cross-correlation modelling technique, Disney et al. found the host galaxy to be excellently matched by an elliptical galaxy model. Their wide-R image confirms the existence of the unresolved companion at \( \approx 2^\prime \) separation from the nucleus and suggests that the environment of 2215−037 resembles a poor cluster. This suggestion is supported by the image of the full WF2 chips which shows \( \geq 10 \) companion objects inside a radius of \( \approx 100 \text{ kpc} \). This quasar was identified as an X-ray source in the ROSAT All-Sky Survey (Yuan et al. 1998).
Figure A1. The radio galaxy 0230−027
Figure A2. The radio galaxy 0307+169
Figure A3. The radio galaxy 1215–033
Figure A4. The radio galaxy 1215+013
Figure A5. The radio galaxy 1330+022
Figure A6. The radio galaxy 1342–016
Figure A7. The radio-loud quasar 1020–103
Figure A8. The radio-loud quasar 1217+023
Figure A9. The radio-loud quasar 2135–147
Figure A10. The radio-loud quasar 2355−082
Figure A11. The radio-quiet quasar 0052+251
Figure A12. The radio-quiet quasar 0204+292
Figure A13. The radio-quiet quasar 1549+203
Figure A14. The radio-quiet quasar 2215−037.
APPENDIX B: LUMINOSITY PROFILES

In this appendix we provide the observed and best-fitting model luminosity profiles for 14 AGN for which the new HST data are reported in this paper. The profiles are followed out to a radius of 10″, which is representative of the typical outer radius used in the image modelling. In each figure the azimuthally averaged data are indicated by circles, the azimuthal average of the best-fitting 2-dimensional model is shown by the solid line, and the dotted line indicates the contribution made to the surface-brightness profile by the point-source component of the model (after convolution with the PSF).

The Radio Galaxy 0230–027

The Radio Galaxy 0307+169

The Radio Galaxy 1215–033
The Radio Loud Quasar 2355–082

The Radio Quiet Quasar 1549+203

The Radio Quiet Quasar 0052+251

The Radio Quiet Quasar 2215–037

The Radio Quiet Quasar 0204+292

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