Decoherence effect in neutrinos produced in microquasar jets

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Abstract. We study the effect of decoherence upon the neutrino spectra produced in microquasar jets. In order to analyse the precession of the polarization vector of neutrinos we have calculated its time evolution by solving the corresponding equations of motion, and by assuming two different scenarios, namely: (i) the mixing between two active neutrinos, and (ii) the mixing between one active and one sterile neutrino. The results of the calculations corresponding to these scenarios show that the onset of decoherence does not depends on the activation of neutrino-neutrino interactions when realistic values of the coupling are used in the calculations. We discuss also the case of neutrinos produced in windy microquasars and compare the results which those obtained with more conventional models of microquasars.

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1 Introduction

The study of neutrino’s related processes in astroparticle physics is a subject of utmost importance since it is strongly connected with crucial aspects of particle physics [1]. The achievements in the field, both theoretically and experimentally are impressive: the values of the neutrino-flavor oscillations parameters have been determined [2–5] and various scenarios for the mass hierarchy have been proposed and constrained experimentally [6]. Lately, the importance of neutrino-neutrino interactions in star evolution has been emphasized [7, 8], particularly in dealing with supernovae’s dynamics. The neutrinos are indeed very peculiar particles, since they can travel enormous distances without being severely affected by local interactions. However, their quantum nature should manifests in phenomena like decoherence [9–11]. As pointed out in ref. [12], the onset of decoherence may affect strongly the density and energy-momentum distribution of neutrinos produced in distant sources.

As it is well known from elementary quantum mechanics, pure states may evolve into mixed states due to interactions with the background [9, 10]. If this is the case with neutrinos produced in supernovae explosions or in other astrophysical events the information about oscillation parameters, masses, etc, may be depending upon the presence of decoherence. In this paper we focus on neutrinos produced from various reactions which take place in a microquasar. By modelling their spectra and initial densities we are able to follow their evolution in time and determine conditions for the appearance of decoherence. We have followed the formalism of [13, 14] and adapted it to calculate the pattern of decoherence in the time evolution of microquasar’s neutrinos. We have found a dependence of the decoherence pattern with the mixing scheme of the neutrinos regardless of the inclusion of neutrino-neutrino interactions. As shown in the calculations, the smallness of the realistic values of
the strength of the interactions allows to neglect them. We have considered two cases for the neutrino mixing, that is a) the mixing between active neutrinos and b) the mixing between active and sterile neutrinos.

The work is organized as follows. In section 2 we present the formalism which we have developed to calculate the time evolution of the neutrino spectra and in section 3 we used it to obtain the neutrino spectrum in a microquasar jet. The results of the calculations are presented in section 4. Our conclusions are drawn in section 5.

2 Formalism

The time evolution of the occupation number of neutrinos is governed by the equation of motion [12]

\[ \dot{\rho}_f = \left[ M^2 c^4 + \sqrt{2} G_F \rho, \rho_f \right], \tag{2.1} \]

where the squared brackets reads for the commutator, the dot represents the time derivative and \( M^2 c^4 \) is the mass-squared matrix in the flavour basis. The quantity \( \rho_f \) is the density matrix in the flavour basis, and \( \rho \) is a matrix whose diagonal terms are the neutrino number densities, \( E \) stands for the neutrino energy and \( G_F \) is the Fermi constant.

The formalism which we are presenting here, as well as the calculations described in the following sections, are limited to the two-flavor scheme for active neutrinos. The use of this limited scheme is justified by the strong mixing between the two lightest neutrino mass eigenstates, which eventually leads to the decoupling from the heavier neutrino mass eigenstate.

The mass-squared matrix in the flavour basis can be obtained by a transformation of the mass-squared matrix in the mass basis \( (m^2 c^4) \) through the unitary mixing matrix \( U \).

That is:

\[
M^2 c^4 = U m^2 c^4 U^\dagger
= \frac{1}{2} \left( m_1^2 + m_2^2 \right) c^4 \left( \begin{array}{cc} 1 & 0 \\ 0 & 1 \end{array} \right) + \frac{1}{2} \sin 2\theta \left( m_2^2 - m_1^2 \right) c^4 \left( \begin{array}{cc} 1 & 0 \\ 0 & -1 \end{array} \right),
\tag{2.2}
\]

In the above equation \( m_1 \) and \( m_2 \) stand for the mass of the neutrino in the mass eigenstate \( \nu_1 \) and \( \nu_2 \), respectively.

Following ref. [12], one can write the mass matrix and the matrix \( \rho_f \) in terms of Pauli matrices. For two-neutrino mass eigenstates they are written

\[
\frac{M^2 c^4}{2 E \hbar} = \frac{1}{2} \text{tr} \left( \frac{m^2 c^4}{2 E \hbar} \right) I + \frac{1}{2} w \vec{B} \cdot \vec{\sigma}, \tag{2.3}
\]

\[
\rho_f = \frac{1}{2} \text{tr} (\rho_f) I + \frac{1}{2} \vec{P}_f \cdot \vec{\sigma}, \tag{2.4}
\]

where \( \vec{P}_f \) is the polarization vector in the flavour basis, \( w = \frac{\delta m^2}{2 E \hbar}, \) and \( \delta m^2 = m_2^2 c^4 - m_1^2 c^4 \) is the mass-squared difference between the mass eigenstates, \( \vec{B} \) is an unitary vector which fixes the orientation of the background. This vector, by simple comparison between the eqs. (2.2) and (2.3), is a column vector with components

\[
\vec{B} = \left( \begin{array}{c} \sin 2\theta \\ 0 \\ -\cos 2\theta \end{array} \right). \tag{2.5}
\]
One can performed a rotation of the coordinate system in the flavour space and choose the direction of $\vec{B}$ in the z axis. This direction does not necessary correspond to the external physical magnetic field. The equation of motion is then re-written as

$$\frac{\partial \vec{P}_w}{\partial t} = (w\vec{B} + \mu \vec{P}) \times \vec{P}_w,$$  

(2.6)

where $\mu$ stands for the strength of the neutrino-neutrino interaction \[12\] and

$$\vec{P} = \int \vec{P}_w dw,$$  

(2.7)

is the total (or global) polarization vector. Then the initial condition for $\vec{P}_w$, is given by

$$\vec{P}_w(0) = \begin{pmatrix} \sin 2\theta \\ \cos 2\theta \end{pmatrix} \cdot g(w),$$  

(2.8)

where $g(w) = A(g_{\mu} - g_e)$, $g_e$ and $g_{\mu}$ are the electron and muon-neutrino spectral functions, $A$ is a normalization constant and $\theta$ is the neutrino mixing angle.

The order parameter that measures coherence is defined as the ratio between the modulus of the perpendicular polarization vector at time $t$ and it at time $t = 0$, that is:

$$R_\theta(t) = \frac{|\vec{P}_\perp(t)|}{|\vec{P}_\perp(0)|},$$  

(2.9)

where $\vec{P}_\perp(t) = \vec{P} - (\vec{P} \cdot \vec{B}) \vec{B}$. The average in angles can be computed as

$$R(t) = \frac{\int R_\theta(t) d\theta}{\int d\theta}.$$  

(2.10)

3 Neutrino spectra

3.1 Gaussian spectrum

Following ref. \[12\] and as toy models we shall use two different Gaussian-like spectra

$$g_{w1}(w) = \frac{1}{2\sqrt{2}\pi}e^{-(w-5)^2/2} + \frac{1}{2\sqrt{2}\pi}e^{-(w+5)^2/2},$$

$$g_{w2}(w) = \frac{1}{2\sqrt{2}\pi}e^{-(w-1)^2/2} + \frac{1}{2\sqrt{2}\pi}e^{-(w+1)^2/2},$$  

(3.1)

that is two-non-overlapping Gaussian distributions, $g_{w1}$, and two-overlapping Gaussian functions, $g_{w2}$, respectively.

3.2 Jet’s neutrino spectrum in a microquasar from $p\gamma$ and $pp$ interactions

In order to compute the neutrino spectrum produced in a microquasar, we follow ref. \[15\]. We have assumed a compact object with an accretion disk and a perpendicular jet with a half-opening angle $\xi$. The injection point is located at $z_0$. In table 1 we show the parameters
### Table 1. Microquasar parameters (some of the values where extracted from ref. [15]).

| Parameter                        | Value       |
|----------------------------------|-------------|
| Jet power                        | $5 \times 10^{40} \text{erg s}^{-1}$ |
| Initial jet’s high               | $10^8 \text{cm}$ |
| Lorentz factor                   | 1.25        |
| Relativistic particles           | 0.1         |
| Hadron-to-lepton ratio           | 100         |
| Half opening angle               | 0.087       |

used in the calculation. The differential equation that gives the density of the particles, $N(E)$, in a microquasar jet is

$$
\frac{\partial (N(E)b(E))}{\partial E} + (t^{-1}_{\text{dec}} + t^{-1}_{\text{esc}}) N(E) = Q(E), \tag{3.2}
$$

where $b(E) = -Et^{-1}_{\text{loss}}$, $Q(E)$ is the particle injection, $t^{-1}_{\text{dec}}$ and $t^{-1}_{\text{esc}}$ are the decay and escape rate respectively. The escape rate can be computed as $t^{-1}_{\text{esc}} = c/(z_{\text{max}} - z)$, where $c$ is the speed of light, $z_{\text{max}}$ is the extent of the acceleration region (that is $5 \times 10^8 \text{cm}$), and $z$ is the position. The rate $Et^{-1}_{\text{loss}}$ is the sum of the cooling rates, that is $t^{-1}_{\text{syn}} + t^{-1}_{\text{ad}} + t^{-1}_{pp} + t^{-1}_{p\gamma}$ for protons, $t^{-1}_{\text{syn}} + t^{-1}_{\text{ad}} + t^{-1}_{\pi p} + t^{-1}_{\pi\gamma}$ for pions and $t^{-1}_{\text{syn}} + t^{-1}_{\text{ad}} + t^{-1}_{ic}$ for muons, respectively. The other quantities needed to evaluate the cooling rates are: $t^{-1}_{\text{syn}}$ which is rate of emission of synchrotron radiation, $t^{-1}_{\text{ad}}$ which stands for the adiabatic cooling, $t^{-1}_{pp}$ which is the $pp$ collision rate, $t^{-1}_{p\gamma}$ which gives the rate of the interaction between protons and synchrotron photons. Finally, $t^{-1}_{\pi p}$ and $t^{-1}_{\pi\gamma}$ represent the proton-pion and pion-photon interaction rates and $t^{-1}_{ic}$ is the loss-rate of the inverse Compton interactions [15]. The solution of equation (3.2) is

$$
N(z,E) = \frac{1}{|b(z,E)|} \int_{E}^{E_{p}^{\text{max}}} Q(z,x)e^{-\tau(E,x,z)}dx \tag{3.3}
$$

where $E > 1.2 \text{ GeV}$. The proton maximum energy $E_{p}^{\text{max}} = 5 \times 10^6 \text{ GeV}$ is obtained through the assumption that the acceleration rate is equal to the loss rate at the initial high of the jet and

$$
\tau(E,x,z) = \int_{E}^{x} dy \frac{t^{-1}_{\text{dec}}(y) + t^{-1}_{\text{esc}}(z)}{|b(z,y)|}. \tag{3.4}
$$

### 3.2.1 Proton injection

The proton injection, $Q(z,E)$ of eq. (3.3) is [15]

$$
Q(z,E) = Q_0^p \left( \frac{z_0}{z} \right)^3 \Gamma^{-1} \left( E - \beta \sqrt{E^2 - m_p^2c^4} \cos \theta \right)^{-2} \left( \Gamma - \cos \theta \frac{E \beta}{\sqrt{E^2 - m_p^2c^4}} \right). \tag{3.5}
$$

$Q_0^p$ is a constant to be determined from the luminosity, $\Gamma$ is the Lorentz factor, $\theta$ is the observation angle and $\beta$ is related to $\Gamma$ (see [15] for details). In this case $t^{-1}_{\text{dec}} = 0$. 


3.2.2 Pion injection

The pion injection is 

\[ \dot{Q}_\pi (E) = Q_{pp}^\pi (E) + Q_{p\gamma}^\pi (E), \]

where \( Q_{pp}^\pi (E) \) and \( Q_{p\gamma}^\pi (E) \) are the injection terms resulting from the pion production due to proton-proton and proton-photon interactions, respectively. The proton-proton injection is calculated as

\[ Q_{pp}^\pi (z, E) = n_p(z)c \int_{E}^{E_{\text{max}}} \frac{E_p}{E_{\text{pp}}} N_p(z, E_p) F_\pi \left( \frac{E}{E_p} \right) \sigma_{pp}(E_p) \frac{dE_p}{E_p}. \]

In the previous expression, \( n_p \) is the density of cold particles, \( N_p(z, E_p) \) is the proton density, \( \sigma_{pp} \) is the cross section and \( F_\pi \) is the pion’s distribution produced per \( pp \) collisions (see refs. [15, 16]).

The proton-photon production is given by the expression

\[ Q_{p\gamma}^\pi (z, E) = 5N_p(z, 5E, \theta)\omega_{p\gamma}(z, 5E) \mathcal{N}_\pi(z, 5E), \quad (3.6) \]

where \( \omega_{p\gamma} \) is the collision-frequency and \( \mathcal{N}_\pi \) the mean number of positive and negative pions.

3.2.3 Muon injection

The pion decay produces muons, therefore the muon injection is

\[ Q_{\mu}^{L^-:R^+} (z, E) = \int_{E}^{E_{\text{max}}} t_{\text{dec} \pi}^{-1} (E_\pi) N_\pi(z, E_\pi) \frac{dn_{\pi^- \rightarrow \mu_L}}{dE} \frac{dE_\pi}{E_\pi}, \]

\[ Q_{\mu}^{R^-:L^+} (z, E) = \int_{E}^{E_{\text{max}}} t_{\text{dec} \pi}^{-1} (E_\pi) N_\pi(z, E_\pi) \frac{dn_{\pi^- \rightarrow \mu_R}}{dE} \frac{dE_\pi}{E_\pi}. \quad (3.7) \]

In the previous expression \( t_{\text{dec} \pi}^{-1} \) is the pion decay rate, \( N_\pi(z, E_\pi) \) the density of pions and \( \frac{dn_{\pi^- \rightarrow \mu_L}}{dE} \) and \( \frac{dn_{\pi^- \rightarrow \mu_R}}{dE} \) are the decay rates of left-handed and right-handed muons, respectively [17].

3.2.4 Neutrino injection

The neutrino production due to pion decay can be written as

\[ Q_{\pi \rightarrow \nu} (z, E) = \int_{E}^{E_{\text{max}}} t_{\text{dec} \pi}^{-1} (E_\pi) N_\pi(z, E_\pi) \Theta \left( 1 - \frac{E}{E_\pi} - \frac{m_\mu}{m_\pi} \right)^2 \left( 1 - \frac{m_\mu}{m_\pi} \right)^{-1} \frac{dE_\pi}{E_\pi}, \quad (3.8) \]

The neutrino injection due to the muon decay is

\[ Q_{\mu \rightarrow \nu} (z, E) = \sum_{i=1}^{4} \int_{E}^{E_{\text{max}}} t_{\text{dec} \mu}^{-1} (E_\mu) N_{\mu_i}(z, E_\mu, y) \left( \frac{E}{E_\mu} \right) \frac{dE_\mu}{E_\mu}, \quad (3.9) \]

where \( t_{\text{dec} \mu}^{-1} \) is the muon decay rate, \( N_{\mu_i}(z, E_\mu, y) \) the muon density and \( y(x) \) a polynomial function [15].

The neutrino spectral function can be calculated by performing the integral in the jet volume

\[ g_\mu(E_\nu) = \int dV \frac{Q_{\mu \rightarrow \nu} (z, E) + Q_{\pi \rightarrow \nu} (z, E)}{t_{\text{esc}}^{-1}} \]

\[ g_e(E_\nu) = \int dV \frac{Q_{\mu \rightarrow \nu} (z, E)}{t_{\text{esc}}^{-1}}. \quad (3.10) \]
| Parameter                  | Value                |
|----------------------------|----------------------|
| $M_*$                      | $10 \, R_\odot$      |
| $M_{\text{compact object}}$| $1.4 \, R_\odot$     |
| $R_*$                      | $10 \, R_\odot$      |
| Period                     | $22.8925 \times 10^4 \, \text{s}$ |
| Eccentricity               | 0.72                 |
| Initial orbital phase      | 0.261799             |
| $\rho_0$                   | $10^{-11} \, \text{gr cm}^{-3}$ |
| $v_\infty$                 | $5 \times 10^5 \, \text{cm s}^{-1}$ |
| Initial jet’s high         | $10^7 \, \text{cm}$  |
| Initial jet’s radius       | $10^6 \, \text{cm}$  |
| Proton spectrum power law  | 2.2                  |
| Lorentz factor             | 1.25                 |

Table 2. Windy microquasar parameters (from ref. [13]).

3.3 Jet’s neutrino spectrum in a windy microquasar

In order to calculate the neutrino density as a function of the energy, we have assumed that a binary system formed by one high-mass primary star surrounded by a disk and a compact object describing a Kepler-orbit (see table 2) and followed the analysis presented in ref. [13]. The jet of relativistic particles produced by the compact object is considered to be cone perpendicular to the accretion-disk plane (or orbital plane). We have used the wind velocity model [14]

$$v(r_w) = v_\infty \left(\frac{R_*}{r_w}\right)^{1.2},$$

(3.11)

where $r_w$ is the radial coordinate from the center of the star, $R_*$ is the star radius, $v_\infty$ is the terminal velocity of the wind. The mass density of the wind is obtained from the continuity equation [12].

Following ref. [14] one can write the proton spectrum in the jet frame and the accretion rate due to the wind to obtain the proton flux in the observer frame [13, 14]. These protons interact with target protons of the wind via the reaction

$$p + p \rightarrow p + p + \xi_{\pi^0}(E_p)\pi^0 + \xi_{\pi}(E_P)(\pi^+ + \pi^-),$$

(3.12)

where $\xi_{\pi^0}(E_p)$ and $\xi_{\pi}(E_P)$ are the multiplicities for neutral and charged pions, given by [13]

$$\xi_{\pi^0}(E_p) = 1.1 \left(\frac{E_p}{\text{GeV}}\right)^{1/4},$$

$$\xi_{\pi}(E_p) = \left(\frac{E_p}{\text{GeV}} - 1.22\right)^{1/5}.$$  

(3.13)

The proton ($E_p$) and photon ($E_\gamma$) energies are related by $E_p = 6k^{-1}\xi_{\pi^0}(E_p)E_\gamma$, where $k = 0.5$ is the inelasticity coefficient.
From the energy conservation one can obtain the neutrino intensity produced by pion- and muon-decay \cite{18, 19}

\[
\int_{E_{\gamma}^{\text{min}}}^{E_{\gamma}^{\text{max}}} dE_{\gamma} \frac{dN_{\gamma}}{dE_{\gamma}} E_{\gamma} = \Delta \int_{E_{\nu}^{\text{min}}}^{E_{\nu}^{\text{max}}} dE_{\nu} \frac{dN_{\nu}}{dE_{\nu}} E_{\nu},
\]

where \(E_{\gamma}^{\text{min}}\) and \(E_{\gamma}^{\text{max}}\) are the minimum and maximum energies of photons resulting from hadrons, and \(E_{\nu}^{\text{min}}\) and \(E_{\nu}^{\text{max}}\) are the corresponding minimum and maximum energy of the neutrinos, and \(\Delta = 1\) \cite{19}. The neutrino energy is related to the photon energy by 

\[
E_{\nu} = \frac{k}{12\xi_{\pi}(E_{p})} E_{p}.
\]

The maximum neutrino energy is determined by the maximum energy acquired by the accelerated protons, which is related to the magnetic field \(B\). The magnetic field is calculated by assuming equipartition between the magnetic field energy and the kinetic energy of the jet \cite{13}. The maximum energy of the protons is

\[
E_{p}(\psi) = eR(z_{0})B(\psi, z_{0}).
\]

Note that \(0.5 \text{ GeV} < E_{p}(\psi) < 2.8 \times 10^{4} \text{ GeV}\).

The orbital-phase dependent muon-neutrino spectrum is computed as

\[
G_{\mu}(E_{\nu}, \psi) = \frac{4f_{p}}{m_{p}} \int dV \rho_{w}(\psi; z, \delta, \phi) q_{\gamma}(\psi; 2E_{\nu}, z, \theta) t_{\text{esc}}^{-1}(z),
\]

where \(f_{p} = 0.1\) takes into account particle-rejection from the boundary \cite{21}, \(\rho_{w}(\psi; z, \delta, \phi)\) is the wind mass density, \(q_{\gamma}(\psi; 2E_{\nu}, z, \theta)\) stands for the gamma-ray emissivity \cite{22} and \(t_{\text{esc}}^{-1}(z) = \frac{c}{z_{m} - z}\) is the inverse of the neutrino’s escape time. The integral is performed in the jet’s volume. The muon neutrino spectrum is then calculated as

\[
g_{\mu}(E) = \frac{1}{2\pi} \int_{0}^{2\pi} d\psi G_{\mu}(E_{\nu}, \psi)
\]

which is further normalized when performing the calculations, as described in the next section. The electron-neutrino spectrum can be determined by repeating the same arguments.

### 4 Results

We have considered two different neutrino’s scenario, that is: (i) two-active neutrinos and (ii) one active and one sterile neutrino, to compute the order parameter. The neutrino-mixing parameters for the first scenario are \(\delta m^{2} = 7.62 \times 10^{-5} \text{ eV}^{2}\) and \(\sin^{2} \theta = 0.307\) \cite{23}. In this case, we have computed both the electron-neutrino and muon-neutrino spectral functions produced in the microquasar and used them as initial condition (see eq. (2.8)) to compute the time dependence of the polarization vector and the order parameter. For the active-sterile neutrino mixing parameters we have used \(\delta m^{2} = 1 \text{ eV}^{2}\) and \(\sin^{2} \theta = 0.1\). As initial condition for the sterile neutrino sector we assume that sterile neutrino are not produced in the microquasar jet.
Figure 1. Left figure: order parameter as a function of the time, for two-separated Gaussian spectra. Top figure: order parameter using the active-active mixing angle; bottom figure: mean value of the order parameter. Solid line: $\mu = 0\, \text{s}^{-1}$; dashed line: $\mu = 1\, \text{s}^{-1}$; dotted line: $\mu = 1.8\, \text{s}^{-1}$. Right figure: order parameter as a function of the time, for two-overlapping Gaussian spectra. Top figure: order parameter using the active-active mixing angle; bottom figure: mean value of the order parameter. Solid line: $\mu = 0\, \text{s}^{-1}$; dashed line: $\mu = 2\, \text{s}^{-1}$; dotted line: $\mu = 3\, \text{s}^{-1}$.

4.1 Gaussian spectra

As toy model we have computed the order parameter for two-active neutrino’s Gaussian spectrum. In figure 1 we show the order parameter for the two-Gaussian spectra at the mixing angle (top figure) and its mean value (bottom figure), as a function of the time and for different values of $\mu$. In absence of neutrino-neutrino interactions, that is $\mu = 0$, the order parameter shrinks to zero as well as its mean value. When the neutrino-neutrino interaction is activated $|\vec{P}|$ decreases its value but oscillates around a non-zero value. The larger the value of the interaction the larger is $|\vec{P}|$. The mean value of the order parameter reaches a smaller average value earlier than the one calculated with a fixed mixing angle.

The results for the case of two-overlapping Gaussian as initial condition of the neutrino spectrum, shown at the right inset of figure 1, are similar to the ones obtained in ref. [12] for a Gaussian spectrum.

4.2 Two-active neutrinos

In this section we present the results for the order parameter calculated by using the neutrino spectral function as described in section 3.2 and in section 3.3. The results for the first case are shown in the left inset of figure 2. They are similar to the ones obtained using as initial condition the two-non-overlapping Gaussian functions. The length of the vector $\vec{P}$ is reduced to zero for the non-interacting case and for small values of $\mu$. However, for larger values of the neutrino-neutrino interaction, the polarization vector oscillates towards an asymptotic non-zero value.
The strength of the neutrino-neutrino interaction, \( \mu \), computed by using realistic values of the neutrino densities \( n_\nu \), that is \( \mu = \sqrt{2}G_F n_\nu \) is of the order of \( 2 \times 10^{-23}\)eV (or in units of inverse time it is of the order of \( 10^{-7}\)s\(^{-1} \)), a value which is orders of magnitude smaller than the strength used in the calculations to illustrate the effects of decoherence. It means that for the case of the mixing between active neutrinos in a microquasar the onset of decoherence is independent of the neutrino-neutrino interactions included in the time derivative of the polarization vector \( \vec{P}_w \) of eq. (2.6). Therefore, in a micro-quasar, the realistic evolution of the order parameter seems to be similar to the one computed without including neutrino-neutrino interaction.

The order parameter \( R \) obtained using the neutrino spectra described in section 3.3 is shown in the right inset of figure 2, for active-active neutrino scheme. For a windy microquasar, the realistic value of \( \mu \) is \( 1.1 \times 10^{-9} \)s\(^{-1} \), therefore, the realistic time evolution of \( R \) is the one obtained with \( \mu = 0 \).

The time that neutrinos remain inside the jet can be computed as \( R_{\text{max}}/c \). For the calculation using the formalism describe in section 3.2, this time is \( 1.4 \times 10^{-3} \)s and in windy microquasar is \( 1.7 \times 10^{-4} \)s. Therefore, inside the microquasar’s jet decoherence has not effect in the neutrino distribution.
Figure 3. **Left figure:** order parameter as a function of the time calculated from the emission of neutrinos in a microquasar’s jet and for active-sterile neutrino scheme. Top figure: order parameter using the active-active mixing angle; bottom figure: mean value of the order parameter. Solid line: $\mu = 0 \text{s}^{-1}$; dashed line: $\mu/w_{\text{max}} = 0.1$; dotted line $\mu/w_{\text{max}} = 0.16$, with $w_{\text{max}} = 6.3 \times 10^5 \text{s}^{-1}$ **Right figure:** order parameter as a function of the time calculated from the neutrino spectrum in a windy microquasar and active-sterile neutrino scheme. Top figure: order parameter using the active-active mixing angle; bottom figure: mean value of the order parameter. Solid line: $\mu = 0 \text{s}^{-1}$; dashed line: $\mu/w_{\text{max}} = 0.33$; dotted line $\mu/w_{\text{max}} = 1.25$. with $w_{\text{max}} = 1.6 \times 10^6 \text{s}^{-1}$.

### 4.3 Active-sterile neutrino

For the active-sterile neutrino, the results are shown in figure 3. The results displayed in the left inset of this figure have been obtained by applying the formalism described in section 3.2. As seen from the curves, the length of the vector $\bar{P}$ is reduced to zero for the non-interacting case but it does not vanishes for unrealistically large values of the strength of the neutrino-neutrino interaction. The realistic value of $\mu$ for this case is of the order of $O(10^{-8} \text{s}^{-1})$, thus the results are consistent with null neutrino-neutrino interactions. We observe that the value of the order parameter is reduced to zero in about $2 \times 10^{-4}$ seconds.

For the windy microquasar jet one can see that, for large values of the parameter $\mu$, the length of the polarization vector is reduced by a 20%, as shown in the right inset of figure 3. Meanwhile, for small values of the interaction, the length of the vector $\bar{P}$ is depleted. The neutrino distribution is strongly affected by decoherence, since the time of permanence of neutrinos in the jet is $O(10^{-3})$ s, for microquasar’s jet and $O(10^{-4})$ s, for a windy microquasar, respectively. For this mixing-scheme the value of $\mu$ computed with the realistic neutrino spectrum is of the order of $O(10^{-10} \text{s}^{-1})$, and the polarization vector is reduced to zero within $4 \times 10^{-5}$ seconds.

Some final comments about the onset of decoherence, for the cases which we have considered in this section. For the case of only active neutrinos, if one fixes the neutrino boundaries in the region of the order of $10^8$ cm the time they spend there is of the order of...
$3.3 \times 10^{-3}$ seconds, a time shorter than the time needed to evolve into pointer states (see figure 2). However, if the coupling with sterile neutrinos is activated the onset of decoherence becomes clear from the results (see figure 3). The same behaviour is then expected for the outside region. A more precise determination of the effect requires the comparison between the length of the flavor oscillations and the length of decoherence.

5 Conclusions

In this work we have studied the effect of collective oscillations upon the neutrino spectral function, for neutrinos produced in micro quasar’s jets, by applying the formalism developed in refs. [13, 15]. Using active neutrinos as initial condition for the evolution of the polarization vector $\vec{P}$ we have calculated the order parameter $R_\theta$ as a function of the time.

For the case of neutrino’s Gaussian spectra we have found that the polarization vector reduces its length to zero for small or null neutrino interactions, exhibiting a complete decoherence-pattern. If the neutrino density is large enough the length of the polarization vector becomes smaller than one and it oscillates around a non-zero asymptotic value. For non-overlapping Gaussians the reduction of the order parameter is faster and the neutrino density must increase in order to reduce the effects of decoherence.

The results using the realistic electron- and muon-neutrino spectra produced in a micro quasar’s jet are quite similar, for both formalisms [13, 15], and similar to the ones obtained with the toy (Gaussian) models. The polarization vector is reduced to zero for small neutrino density while for larger values of the neutrino-neutrino interactions, the decoherence is not completed. However, decoherence inside the microquasar’s jet does not have a strong effect upon the polarization vector. The realistic value of $\mu$ is quite small, therefore the time-evolution of the order parameter is not affected by the presence of neutrino-neutrino interactions.

When a sterile neutrino is oscillating with a light active-neutrino (electron-neutrino) the effects of the collective oscillations are noticeable at very small times, since the maximum value for the frequency $w$ is quite large due to the mass difference. In this case, for the electron-neutrino spectral function calculated in section 3.2, the decoherence is almost complete for the mixing angle used, but the mean value of the order parameter is different from zero for large values of the neutrino-neutrino interaction. For the case of a windy microquasar the polarization vector is reduced and this reduction is smaller if the neutrino density is larger. Inside the jet, in this scheme, the decoherence strongly affects the neutrino distribution function. In these cases, the realistic strength of neutrino-neutrino interactions is smaller than the corresponding to the case of the mixing between two-active neutrinos. We have found that decoherence is completed within $10^{-4}$ s for a micro quasar and $4 \times 10^{-5}$ s for a windy micro quasar, respectively.

Finally, the effect of decoherence is present in the neutrino flux produced in microquasar jets, with or without sterile neutrino. This effect is noticeable for small values of the neutrino density. For active-sterile neutrino mixing the effect becomes noticeable at earlier times than for the mixing between active neutrinos.

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