PDF hosted at the Radboud Repository of the Radboud University Nijmegen

The following full text is a preprint version which may differ from the publisher's version.

For additional information about this publication click this link.
http://hdl.handle.net/2066/143843

Please be advised that this information was generated on 2020-03-11 and may be subject to change.
The evolution of a jet ejection of the ultraluminous X-ray source Holmberg II X-1

D. Cseh, J. C. A. Miller-Jones, P. G. Jonker, F. Grisé, Z. Paragi, S. Corbel, H. Falcke, S. Frey, P. Kaaret, E. Körding

1 Department of Astrophysics/IMAPP, Radboud University Nijmegen, P.O. Box 9010, 6500 GL Nijmegen, The Netherlands
2 International Centre for Radio Astronomy Research - Curtin University, GPO Box U1987, Perth, WA 6845, Australia
3 SRON, Netherlands Institute for Space Research, Sorbonnelaan 2, 3584 CA Utrecht, The Netherlands
4 Observatoire astronomique de Strasbourg, Université de Strasbourg, CNRS, UMR 7550, 11 rue de l’Université, F-67000 Strasbourg, France
5 Joint Institute for VLBI in Europe, Postbus 2, 7990 AA Dwingeloo, The Netherlands
6 Laboratoire AIM (CEA/IRFU-CNRS/INSU-Université Paris Diderot), CEA/DSM/IRFU/SAp, F-91191 Gif-sur-Yvette, France
7 Station de Radioastronomie de Nançay, Observatoire de Paris, CNRS/INSU, USR 704 - Univ. Orléans, OSUC, 18330 Nançay, France
8 FOMI Satellite Geodetic Observatory, P.O. Box 585, H-1592 Budapest, Hungary
9 Department of Physics and Astronomy, University of Iowa, Iowa City, 52240, USA

ABSTRACT
We present quasi-simultaneous, multi-epoch radio and X-ray measurements of Holmberg II X-1 using the European VLBI Network (EVN), the Karl G. Jansky Very Large Array (VLA), and the Chandra and Swift X-ray telescopes. The X-ray data show apparently hard spectra with steady X-ray luminosities 4 months apart from each other. In the high-resolution EVN radio observations, we have detected an extended milli-arcsecond scale source with unboosted radio emission. The source emits non-thermal, likely optically thin synchrotron emission and its morphology is consistent with a jet ejection. The 9-GHz VLA data show an arcsecond-scale triple structure of Holmberg II X-1 similar to that seen at lower frequencies. However, we find that the central ejection has faded by at least a factor of 7.3 over 1.5 years. We estimate the dynamical age of the ejection to be higher than 2.1 years. We show that such a rapid cooling can be explained with simple adiabatic expansion losses. These properties of Holmberg II X-1 imply that ULX radio bubbles may be inflated by ejecta instead of self-absorbed compact jets.

Key words: accretion, accretion discs – black hole physics – X-rays: binaries

1 INTRODUCTION
Ultraluminous X-ray sources (ULXs) host accreting compact objects, whose luminosity (L_X > 3 \times 10^{39} \text{ erg/s}) exceeds the Eddington limit of a 20-M_\odot black hole (BH). ULXs are a mixed bag of sources. Some may host super-Eddington neutron stars (Bachetti et al. 2014) or stellar-mass (M = 3–20 M_\odot) BHs (Liu et al. 2013; Motch et al. 2014). Some of them are likely to be intermediate-mass M = 10^2–10^5 M_\odot BHs (Farrell et al. 2009, Feng & Kaaret 2010, Pasham et al. 2014, Mezcua et al. 2015) and may also host recoiling supermassive BHs (Jonker et al. 2010).

Given the tentative observational evidence for the anti-correlation between BH mass and host galaxy metallicity (Crowther et al. 2010) and that some ULX progenitor stars could potentially have been massive enough (e.g. Grisé et al. 2012, Pontanen et al. 2013) to collapse directly to a BH, ULXs might host massive stellar-mass (20-100 M_\odot) BHs (Fryer 1999, Mapelli et al. 2009, Zampieri & Roberts 2009, Belczynski et al. 2010, Mapelli & Bressan 2013, Bachetti et al. 2013). In fact, Holmberg II X-1 lies in a relatively metal-poor dwarf irregular galaxy (Prestwich et al. 2013, Brorby et al. 2013). Such a mass range has been shown to fit its observed X-ray and radio properties (Goad et al. 2008, Cseh et al. 2014) and may also be valid for IC 342 X-1 (Marlowe et al. 2014).

To date, relatively few ULX radio counterparts have been studied, and these few show a relatively diverse nature. The first radio bubble, associated with NGC 5408 X-1, was discovered about a decade ago (Kaaret et al. 2003) although one other, associated with NGC 6946 X-1, was detected earlier but considered as a super- or hypernova.
remnant (van Dyk et al. 1994). Following the discoveries of more radio bubbles (e.g., Cseh et al. 2012 and references therein), transient jets were also caught in two non-persistent ULXs, which in both cases were associated with X-ray outbursts of the central compact object (Webb et al. 2012; Middleton et al. 2013). Furthermore, in a few other cases, apparently steady radio emission has been detected on milli-arcsecond scales (Mezcua & Lobanov 2014; Mezcua et al. 2013, 2014) and interpreted as self-absorbed radio cores, which are typically seen during radiatively inefficient accretion states of Galactic BH X-ray binaries (e.g., Fender et al. 2009).

Focusing on the large-scale environments of ULXs, jets have been argued to be a plausible supply mechanism responsible for inflating the observed radio bubbles (and shock-ionized optical bubbles) (e.g., Cseh et al. 2012). On the other hand, direct evidence of jets in Holmberg II X-1 (Cseh et al. 2014) has been found, despite an environment that appears to be dominated by photoionisation (Lohman et al. 2003; Egorov et al. 2013). This may indicate that shock ionization is not a necessary consequence of jets. Shock-ionized optical bubbles (e.g., Pakull & Mirioni 2002; Pakull & Grindlay 2008) could equally be inflated by powerful winds. These winds are thought to originate from a supercritical disk that may also give rise to geometrically collimated X-rays (Pontan et al. 2007), although calorimetry of photoionized optical bubbles argues against strongly beamed X-ray emission (e.g., Pakull & Mirioni 2002; Kaaret & Corbel 2009). Moreover, direct observational evidence for super-Eddington winds in ULXs is scarce (Walton et al. 2012, 2013; Middleton et al. 2014). Interestingly, a similar group of objects, such as S26 in NGC 7793 and MQ1 in M83, which do not appear to have luminous central X-ray sources, can nonetheless have more energetic optical bubbles (e.g., Pakull et al. 2014; Soria et al. 2014) than classical ULXs. The detailed structure in Ho II X-1 coupled with a high X-ray flux suggests that the jet morphology indicates a recurrent activity (see below). Also, the above mass range might be supported by the longer activity of the source as compared to Galactic microquasars (Cseh et al. 2014).

How these massive stellar-mass BHs evolve remains unclear. Early BH growth is often hypothesised to proceed via (intermittent) super-Eddington accretion (e.g., Volonteri & Silk 2014; Pacucci & Ferrara 2013). To investigate if some ULXs and their observable properties can act as a proxy for super-Eddington BH growth in the early Universe, we present quasi-simultaneous and multi-epoch observations of the ULX Holmberg II X-1 (hereafter Ho II X-1), which are summarised in Table 1. Our primary aim is to study the nature and evolution of the previously-detected central radio source (Cseh et al. 2014), and to investigate the link between the radio emission and the X-ray behaviour. In Sec. 2 we describe our observational results. In Sec. 2.1 we first discuss the X-ray spectra and the high-resolution radio properties of Ho II X-1, followed by our interpretation of the temporal evolution of the radio ejecta in Ho II X-1. Finally, we discuss the physical properties of the jet ejecta.

## 2 OBSERVATIONS, ANALYSIS, AND RESULTS

### 2.1 Joint Swift/XRT and EVN observations

#### 2.1.1 Swift/XRT observations

The Swift X-ray telescope (XRT) observed Ho II X-1 (PI: Grisé) in photon counting (PC) mode on 17, 20, and 22 Jan 2014 starting from 13:36:59 UT, 00:48:59 UT, and 02:27:58 UT, for useful exposures of 5.71 ks, 5.88 ks, and 8.96 ks, respectively. We retrieved level 2 event files that were sub-

| Table 1. Summary of the observations |
|-------------------------------------|
| Inst. | Mode or config. | Frequency or energy band | On-source time | Observation date |
| EVN   | –              | 1.6 GHz                | 12.5 h         | 14 Jan 14       |
| Swift | PC             | 0.3 – 10 keV           | 20.5 ks        | 17-22 Jan 14    |
| EVN   | –              | 5 GHz                  | 12.5 h         | 28 Apr 14       |
| Chandra | 1/8            | 0.3 – 10 keV           | 11.8 ks        | 25 May 14       |
| VLA   | A              | 8 – 10 GHz             | 171 min        | 25 May 14       |
jected to standard screening methods\textsuperscript{1}. The source spectrum was extracted using an aperture of 20 pixels, which corresponds to 90% of the PSF at 1.5 keV. The background was extracted using an aperture of radius 40 pixels, centred on a J2000 location of RA=08:19:51.437 and Dec=+70:47:46.29. The auxiliary response file (ARF) was created with XRTMKARF, which accounts for PSF correction.

After rebinning the spectrum to have a minimum of 30 counts per bin, we fit the X-ray spectrum using XSPEC v12.8.2 (Arnaud 1996) and a response matrix of swxpc0to12s6. We fit the spectrum with an absorbed power law model (TBABS*TBABS*POWERLAW) and froze the first hydrogen absorption column density to the Galactic value of $N_H = 3.4 \times 10^{20}$ cm$^{-2}$ (Dickey & Lockman 1990). We found a good fit of $\chi^2$/DoF = 0.99 for 80 degrees of freedom (DoF), a photon index of $\Gamma = 1.0$, and an equivalent hydrogen absorption column density of $N_H = (5 \pm 2) \times 10^{20}$ cm$^{-2}$, and an unabsorbed flux of $(6.6 \pm 0.4) \times 10^{-12}$ erg cm$^{-2}$ s$^{-1}$ in the 0.3–10 keV energy range. The fitted power-law photon index of $\Gamma = 1.91$ is consistent with a hard X-ray spectrum. The flux corresponds to a luminosity of $(9 \pm 3) \times 10^{39}$ erg s$^{-1}$ at the source distance of 3.39 Mpc (Karachentsev et al. 2002). We fit other models to the X-ray data. First a disk blackbody model, that gave a poor fit with a reduced $\chi^2$/DoF=1.94, may indicate an over-constrained model. The first hydrogen absorption column density was frozen to the Galactic value. The fit did not require any additional $N_H$ on top of the Galactic one. The fit resulted in inner disc temperatures of $T_{in,1} = 0.32 \pm 0.05$ keV and $T_{in,2} = 1.7^{+0.3}_{-0.2}$ keV. The hotter component agrees within the errors with the result of Walton et al. (2013) and the cooler component is inconsistent within the errors; it is somewhat hotter than found by Walton et al. (2013).

Motivated by the best-fit model of Walton et al. (2013), we also fit the Swift spectrum with two disc blackbody models (TBABS*TBABS*(DISKBB+DISKBB)). This fit resulted in a reduced $\chi^2$/DoF=0.94, that may indicate an over-constrained model. The first hydrogen absorption column density was frozen to the Galactic value. The fit did not require any additional $N_H$ on top of the Galactic one. The fit resulted in inner disc temperatures of $T_{in,1} = 0.32 \pm 0.05$ keV and $T_{in,2} = 1.7^{+0.3}_{-0.2}$ keV. The hotter component agrees within the errors with the result of Walton et al. (2013) and the cooler component is inconsistent within the errors; it is somewhat hotter than found by Walton et al. (2013). The observations were carried out in standard phase-referencing mode, using J0841+7053 as the phase-calibrator. The observations cycled continuously between the target and the reference source, spending 3.5 min on the target in every 5-min block. The data were calibrated in AIPS v31Dec12 (e.g. Greisen 2003) following standard procedures (Diamond 1993), and imaging was performed using DIFMAP (Shepherd et al. 1994). Antenna-based gain correction factors were obtained from the phase calibrator after a global amplitude self-calibration step, which were fed back to AIPS and applied to the data. These corrections were typically less than 10% at 1.6 GHz and less than 20% at 5 GHz. Fringe-fitting was then repeated taking into account the clean component model of the calibrator.

At 5 GHz, we reached an rms noise level of 12 $\mu$Jy beam$^{-1}$, the Gaussian restoring beam size was 2.7 mas $\times$ 2.1 mas, and we did not detect Ho II X-1. We set a 3-$\sigma$ upper limit of 36 $\mu$Jy on the flux density of the source.

At 1.6 GHz, the rms noise level was $\sim$5 $\mu$Jy beam$^{-1}$ and the Gaussian restoring beam size was 12 mas $\times$ 9.6 mas in PA 32$^\circ$. Ho II X-1 was detected at a signal-to-noise ratio of $>10$ (Fig. 1), with a peak brightness of 53 $\mu$Jy at a position of (J2000) RA=08:19:28.935 and Dec=+70:42:18.9966 (accurate to within 1 mas). To check the compactness of the source we fit the brightness distribution with a Gaussian in both the image and the visibility plane. Using the AIPS task JMFIT for the image-plane fit, we found a fitted peak

\textsuperscript{1} http://heasarc.nasa.gov/docs/swift/analysis/
\textsuperscript{2} http://heasarc.gsfc.nasa.gov/docs/heasarc/caldb/-data/swift/xrt/index.html
brightness of $49 \pm 6 \mu\text{Jy beam}^{-1}$ and an integrated flux density of $84 \pm 14 \mu\text{Jy}$. The integrated flux density is considerably higher than the peak brightness, likely indicating a resolved source.

A fit in the visibility plane using a point-source model in DIFMAP resulted in a total flux density of $S_\nu = 65 \pm 6 \mu\text{Jy}$. If we fit the source with a circular Gaussian model, we find a total flux density of $S_\nu = 88 \pm 11 \mu\text{Jy}$ and a full width at half maximum (FWHM) size of $16.4 \pm 1.7$ mas. A comparison of these flux densities also indicates that the source is likely resolved. The obtained size of $16.4 \pm 1.7$ mas exceeds the restoring beam size of $12 \times 9.6$ mas. Furthermore, we estimate a resolution limit of $3\times \mu\text{s}$ that is well sampled by the image cell size of 1 mas.

To test at what angular resolution the source starts to be resolved, we produced an image with a restricted $(u,v)$ range of $0–4 \text{ M}\lambda$. The resulting image had a beam size of $33 \times 30$ mas. The source had a peak brightness of $67 \mu\text{Jy beam}^{-1}$, and a fit to a circular Gaussian model gave a total flux density of $95\pm 10 \mu\text{Jy}$ with a FWHM of $20 \pm 1.6$ mas. There is an apparent increase in all of the fit parameters, although they are consistent with the above values within the errors. We conclude that the source is resolved on VLB1 scales. We set a lower limit to its angular size of $16.4 \pm 1.7$ mas that corresponds to a projected size of $0.26 \pm 0.03$ pc at a distance of Ho II X-1. A short summary is shown in Table 2. Should the true size be larger than $16.4$ mas, we estimate an upper limit of $\sim 60$ mas from our highest resolution VLA images.

### 2.2 Joint Chandra and VLA observations

#### 2.2.1 Chandra observations

We observed Ho II X-1 with the Chandra X-ray Observatory under ObsID 15771 (PI: Jonker) using the Advanced CCD Imaging Spectroscopic array (ACIS) S3 chip in FAINT mode and with 1/8 subarray. The observation was carried out on 25 May 2014, started at 03:21:52 UT, and had a useful exposure of 11.85 ks. There were no strong background flares. The data were processed in a standard way and we used the CIAO version 4.6 and the CALDB version 4.6.1 to further process the level 2 event files.

First, we checked for background flares following standard procedure. We removed strong sources and then checked the extracted background light curve and we found no strong flares during the observation. However, Ho II X-1 is a relatively bright X-ray source that caused a read-out streak, which we therefore removed as described in the Chandra science thread page.

Using WAVDETECT, we detected Ho II X-1 with a total of 6406 photons, and extracted its spectrum in the $0.3–10$ keV range using SPECEXTRACT, with a circular aperture of radius 3 arcsec.

Given that we employed 1/8 subarray, providing a frame time of 0.4 sec, the expected pileup fraction of Ho II X-1 is $\sim 5\%$. To fit the spectrum we used the Sherpa package and its built-in pile-up model, following the Sherpa thread page. First, we grouped our spectrum to have 30 counts per bin and then subtracted the background. We used an absorbed power-law model ($\times$TBABS*$\times$TBABS*$\times$POWERLAW), setting the initial value of the hydrogen absorption column density to the Galactic one of $N_H = 3.4 \times 10^{20}$ cm$^{-2}$. We used the 'Neldermead' optimisation method to find a global solution that resulted in a reduced $\chi^2$/DoF of 1.14, a best-fit value of the photon index of $\Gamma = 1.88 \pm 0.13$, and a hydrogen absorption column density of $N_H = (5\pm 2) \times 10^{20}$ cm$^{-2}$. The resulting pile-up fraction was $\sim 6\%$, in agreement with our estimate above.

To determine the flux we fit the spectra without a pile-up model and froze the power-law index to $\Gamma = 1.88$. We found an unabsorbed flux of $(6.5 \pm 0.6) \times 10^{-12}$ erg cm$^{-2}$ s$^{-1}$ in the $0.3–10$ keV range, corresponding to $L_X = (9.0 \pm 0.8) \times 10^{39}$ erg s$^{-1}$ at the source distance of 3.39 Mpc. We also attempted to fit the same X-ray models to the Chandra spectrum as in Sec. 2.2.1. The spectrum is well fit by a power-law and does not require any additional component. The power law fit parameters and the X-ray luminosity are consistent with being the same for the Swift and the Chandra observations, which were separated by $\sim 4$ months.

Following Walton et al. (2015), we fit the Chandra spectrum with two disc blackbody models in XSPEC (PILEUP*$\times$TBABS*$\times$TBABS* (DISKBB+DISKBB)). We also fit a disc blackbody plus a p-free disk blackbody model (PILEUP*$\times$TBABS*$\times$TBABS* (DISKBB+DISKBB)). These fits resulted in a reduced $\chi^2$/DoF of 1.10 and $\chi^2$/DoF of 1.09, respectively. These reduced $\chi^2$ values may indicate that employing two disks instead of a power law does not significantly improve the fit. The first hydrogen absorption column densities were frozen to the Galactic value. In both of the cases, the fit did not require any additional $N_H$ on top of the Galactic one. The first fit resulted in inner disc temperatures of $T_{in,1} = 0.31 \pm 0.04$ keV and $T_{in,2} = 1.6^{+0.3}_{-0.2}$ keV. The second fit resulted in $T_{in,1} = 0.28 \pm 0.07$ keV, $T_{in,2} = 2.1^{+1.3}_{-0.2}$ keV, and $p = 0.6^{+0.3}_{-0.1}$. The corresponding fitted temperatures between the first and the second fit agree within the errors. Furthermore, given that the value of $p$ can be consistent with 0.75, with the standard disc, we cannot differentiate between a standard or a p-free disc model.

The fitted parameters of the first fit are consistent with being the same for the Swift and the Chandra observations. However, the cooler component is somewhat hotter than found by Walton et al. (2015). On the other hand, if a disk blackbody plus p-free disk is employed, both, the cooler and hotter component agrees within the errors with the result of Walton et al. (2015).

#### 2.2.2 VLA observations

The Karl G. Jansky Very Large Array (VLA) observed Ho II X-1 simultaneously with Chandra on 25 May 2014. The observations were carried out under project code SF0330 and ran from 03:44:50 UT for a total of 3.5 h. The VLA was in its most-extended A-configuration, and the correlator integration time was 2 s. Two 1-GHz wide basebands covered the 8–10 GHz range (X band) with a total of 16 spectral

---

3 The error on this and the subsequent flux densities includes a $5\%$ systematic uncertainty.

4 http://cxc.harvard.edu/ciao/threads/filter_lctrv/

5 http://cxc.harvard.edu/ciao/threads/acisreadcorr/

6 http://cxc.harvard.edu/toolkit/pimms.jsp

7 http://cxc.harvard.edu/sherpa/threads/pileup/index.html
Table 2. Summary of results on the central component

| Instrument | Band      | Flux density [μJy] | θ [mas] | $T_b$ [K] |
|------------|-----------|-------------------|---------|-----------|
| EVN        | 1.6 GHz   | 88±11             | 16.4±1.7| 1.5×10^5  |
| EVN        | 5 GHz     | <36               | –       | –         |
| VLA        | 8 – 10 GHz| 12±5              | <235    | –         |

| Instrument | Band      | Unabsorbed flux [erg cm^{-2} s^{-1}] | Γ | $N_H$ [cm^{-2}] |
|------------|-----------|-------------------------------------|---|----------------|
| Swift      | 0.3–10 keV| (6.6±0.4)×10^{-12}                  | 1.91±0.08 | (5 ± 2)×10^{20} |
| Chandra    | 0.3–10 keV| (6.5±0.6)×10^{-12}                  | 1.88±0.13 | (5 ± 2)×10^{20} |

The results of the Gaussian fits to the central radio component and the fit of an absorbed power-law model to the X-ray source. For the 1.6-GHz VLBI data, the brightness temperature ($T_b$) value is also provided.

Figure 2. The VLA image of Ho II X-1 in X band, using Briggs robust=1 weighting. The lowest contours start at three times the rms noise level, at ±6.6 μJy beam^{-1}. Contours are increased by factors of $\sqrt{2}$. The peak brightness is 15 μJy beam^{-1}. The Gaussian restoring beam size is 0.41″ × 0.22″ at a major axis position angle of 75° E of N. The background colour image is arbitrarily scaled.

3 DISCUSSION

3.1 Is Ho II X-1 in the canonical X-ray hard state?

It is common to fit ULX spectra with a two-component phenomenological model such as a power law and disk blackbody. On the other hand, high-quality XMM-Newton spectra of ULXs show a soft excess and a luminosity-dependent spectral break (Stobbart et al. 2006; Gladstone et al. 2009), and this is also the case for Ho II X-1 (Kajava et al. 2012), that may indicate super-Eddington accretion. On the other hand, Caballero-García & Fabian (2010) suggested that such a spectrum might be reflection-dominated and that the dominant contribution to the total luminosity is coming from the power-law component, potentially implying the sources to be sub-Eddington. However, in the case of Ho II X-1, this model resulted in a fitted photon index that might be too steep for a canonical hard
state (Caballero-García & Fabian 2010). Furthermore, NuSTAR observations have confirmed the presence of a high-energy cutoff for a number of ULXs (Bachetti et al. 2013; Walton et al. 2014; Rana et al. 2015), which in those cases likely rules out a Compton hump arising from a reflection component in a canonical hard X-ray state. This has recently been shown specifically for Ho II X-1 (Walton et al. 2013).

Thus, despite the fact that the current spectra are well fit by a single power law, we argue that Ho II X-1 is unlikely to be in the canonical hard state. Additionally, there is no hint of a self-absorbed compact jet in Ho II X-1 (see later), as is typically seen in this state. It is also unlikely that Ho II X-1 is in a non-standard, efficient hard state (Rushton et al. 2011; Coriat et al. 2011) since the X-rays are steady and decorrelated from the radio emission. Also, our X-ray spectra may as well be fit with two disk blackbodies. In this case, it would be straightforward to exclude Ho II X-1 being in a canonical hard state, and such a fit possibly indicates a non-standard high state or a super-Eddington accretion state (Walton et al. 2013).

In the following we investigate the temporal behaviour of the X-ray spectra. Grisé et al. (2010) presented multiple Swift/XRT observations of Ho II X-1, finding the source to remain in a soft spectral state for over 4 months in 2009/2010, in contrast to the hard spectra observed in 2006, which persisted for up to a month. The hard spectra from 2006 could also be fit with a disk blackbody and a power law model. This is contrary to our spectra where this combination of models resulted in a 0 keV inner disk temperature. On the other hand, the power-law components are identical within uncertainties between these Swift data sets, so a variable soft component is likely to be present. Also, comparing our Swift and Chandra spectra, the power-law components are consistent with being the same even though they are separated by 4 months. Based on our observations, we conclude that the (apparently) hard spectra are in general steady, as compared to the soft spectra, which, beyond a possibly variable soft component, are associated with irregular flares (Grisé et al. 2010). The potential link between these flares and the radio activity needs further investigation.

3.3 Time-dependent radio behaviour

Assuming a non-variable source, the spectral index of the central component between 5.24 and 9.0 GHz is $\alpha_C = -4.7 \pm 0.5$, which is significantly steeper than the previous estimate of $\alpha = -0.8 \pm 0.2$ (Cseh et al. 2014). In fact, such a steep spectrum is unphysical unless the synchrotron spectrum has a cutoff at the observed frequency due to, e.g., synchrotron ageing. We investigate this option in the following section.

Since the two VLA measurements were taken about 1.5 yr apart, we therefore consider time variability. Taking the 5.24-GHz VLA peak brightness of 152 $\mu$Jy/beam and a spectral index of $\alpha = -0.8 \pm 0.2$, the expected peak brightness of a point source at 9 GHz is $99 \pm 11$ $\mu$Jy/beam. The detected X-band peak brightness of $\sim 12$ $\mu$Jy/beam is at least a factor of 7.3 lower than this (a difference of $\sim 8\sigma$). Given that resolution effects were excluded (Sec. 2.2.2) and that the central source is point-like at both C and X band, this indicates that the flux density decreased with time.

3.4 The jet properties of Ho II X-1

The central component is due to optically thin synchrotron emission, consistent with a compact ejection. Since it was found to be a factor of $\sim 3$ brighter than the outer components, this was interpreted as a renewed radio activity in the core (Cseh et al. 2014). The current VLA observation, obtained 1.5 yr later, also shows the same morphology. However, the brightness of the central component is comparable to or less than that of the outer components. Therefore, the
central component has cooled much faster than the outer components. This is consistent with the deduced spectral indices in Sec. 2.2.2. The central component has apparently faded by at least a factor of 7.3, is resolved on VLBI scales, and is not affected by Doppler boosting.

The variation in the flux density over the course of ~1.5 yr is likely a result of expansion. On the other hand, without a light curve, and with a VLBI component that is not (yet) resolved into a two-sided jet, the evolutionary stage of the ejection is unclear. Here, we assume a spreading jet and a Mach number similar to that of the large-scale jet (Cseh et al. 2014). Then a lateral expansion velocity of 0.2c is expected for a speed of c along the jet axis. This maximum expansion speed would result in a maximum radial increase of ~8 mas per year and may limit the age of the ejection to ≥2.1 years. On the other hand, the bulk speed is unknown and based on the geometry it could be as slow as 0.17c. Considering such a mildly relativistic bulk speed, the radial expansion rate would be ~1.3 mas/yr, implying an increased ejection age of ~13–46 yr for our measured size of 16.4–60 mas.

In the following, we deduce an approximate adiabatic timescale ($t_{\text{ad}}$) and compare it with the synchrotron cooling timescale ($\tau$). By assuming a fixed expansion speed of $v_{\text{exp}} = 0.2c$, the adiabatic timescale can be approximated from the apparent physical size ($r$) of the ejection using $t_{\text{ad}} \approx \frac{\tau}{2} \approx \frac{r}{v_{\text{exp}}}$. For an inner structure of size 16.4–60 mas, this gives $t_{\text{ad}} \approx 3.2–11.7$ yr. Next, we estimate the synchrotron lifetime of relativistic electrons at an observed frequency of 9 GHz, using $\tau \approx 2.6 \times 10^3 \nu^{-1/2} B^{3/2}$ (Longair 1994), where the units are in Hz, mG, and yr. We estimate the magnetic field strength from minimum energy arguments, assuming equipartition between the energy in relativistic particles and that in the magnetic field. Using $B = 1.8 \times 10^7 \left( \frac{\nu \xi_0}{\nu} \right)^{2/7} \nu^{1/7}$ mG (e.g. Longair 1994), and substituting in the VLBI parameters of $\nu = 1.6 \times 10^9$ Hz, $V = 3 \times 10^{27}$ m$^3$, and $L_{\nu} = 1.21 \times 10^{47}$ WHz$^{-1}$, we find $B \approx 0.77 \nu^{2/7}$ mG, where $\xi_0$ is the ratio of energy in protons to that in relativistic electrons. Here we consider values of $\xi_0 = 1–100$, and find $\tau \approx (1.8–13) \times 10^3$ yr$^2$. A comparison of these timescales indicates that the ejection is dominated by adiabatic losses, which results in the observed rapid flux decay. The outer components are bright because they are likely the working surfaces where the jet ejecta interact with the surroundings, allowing the conversion of kinetic energy to radiation. Finally, we note that with an average jet power of $Q_j \approx 2.1 \times 10^{39}$ erg s$^{-1}$, the surrounding radio bubble might have been inflated in just ~390 yr (Cseh et al. 2014).

### 3.5 Summary and implications

We have conducted quasi-contemporaneous X-ray and radio observations of Ho II X-1. The X-ray observations show steady, apparently hard spectra with properties that are identical to the power-law component observed in spectra from 2006 (Grisé et al. 2010). By contrast, other observations in 2009/2010 have shown soft spectra and irregular flaring activity (Grisé et al. 2011). Given the steady, apparently hard emission, the X-ray flux variability may well arise purely from the soft component. The VLBI data reveal an extended radio source with a size of 0.26 pc. Based on its variability and brightness temperature, we find that it is consistent with non-thermal emission and that it is not affected by Doppler boosting. This is consistent with a scenario where the radio emission is due to a jet ejection event. Assuming a laterally-expanding jet, we estimate the dynamical age of the ejection to be ≥2.1 yr. Whether an ejection event could be correlated with the soft or flaring nature of the X-ray spectra remains to be investigated. On the other hand, we argue that the hard X-ray spectra are unlikely to correspond to a canonical hard state. This conclusion is also supported by the lack of a classical flat-spectrum, partially self-absorbed compact radio jet. Also, the lack of any correlation between the X-rays and the radio emission may also rule out radiatively efficient hard states, i.e. the ‘outlier branch’ on the X-ray/radio correlation of BH X-ray binaries (e.g. Coriat et al. 2011). Using the VLA we have detected a triple structure at 8–10 GHz that corresponds well to the morphology seen at lower frequencies. However, the central component has faded by at least a factor of 7.3 in just 1.5 years. We explain the observed cooling with adiabatic expansion losses that dominate over synchrotron cooling, regardless of the jet composition (baryonic or leptonic).

Cseh et al. (2012) reported unresolved radio emission in the vicinity of the ULX IC 342 X-1 and concluded that it was inconsistent with a self-absorbed compact jet, and could instead be a clump of emission from the radio nebula. On the other hand, Marlowe et al. (2014) argued that this unresolved emission could be consistent with an expanded ejection similar to that in Ho II X-1. By comparison, it is indeed plausible that such a radio component in IC342 X-1 could have faded on a timescale of a few years. If that is the case, then the quoted size limit on the ejection (Marlowe et al. 2014) of IC 342 X-1 might be too large, because the spatial variation would be small even with highly relativistic expansion speeds.

In light of our results on Ho II X-1, it is possible that other ULX radio bubbles could also have been inflated by transient ejecta. This could imply that it is not necessary to invoke self-absorbed compact jets to inflate ULX bubbles. Moreover, the self-absorbed compact jet in Cyg X-1 has been argued to be incapable of inflating its bubble (Heinz 2008) due to the mismatch between the energetics of the jet and the bubble. On the other hand, expanding ejecta and disk winds, which are potentially triggered by high accretion-rate events, could serve as a more plausible explanation in the above cases, as they might indeed be capable of carrying the required energy.

Ho II X-1 shows a collimated jet structure that is coupled with a high average X-ray luminosity whose ionizing rate is comparable to the average jet power. This may be unexpected given that AGNs with high accretion rates show strong winds but no powerful jets (e.g. Gan et al. 2014, Tombesi et al. 2015). Similarly, in the context of regular BH X-ray binaries the jet is also suppressed in disc-dominated states, and Meier (1996) argues that a super-Eddington wind could be a secondary jet-suppression mechanism. Nonethe-
less, Ho II X-1 does show discrete jet ejecta. This may imply that if the current X-ray luminosity is super-Eddington, then Ho II X-1 has to switch between sub- and super-Eddington rates to form jet ejecta. Alternatively, it could be that the source is not super-Eddington, but is rather in a persistent high, intermediate state where jet ejecta are thought to be released whenever crossing the so-called ‘jet line’.

X-ray monitoring of the source with triggered radio observations could differentiate between any of the cases and test if X-ray flaring traces a jet ejection event.

ACKNOWLEDGMENTS

DC thanks S. Thoudam for useful discussions. The National Radio Astronomy Observatory is a facility of the National Science Foundation operated under cooperative agreement by Associated Universities, Inc. The scientific results reported in this article are based in part on observations made by the Chandra X-ray Observatory. The European VLBI Network is a joint facility of European, Chinese, South African, and other radio astronomy institutes funded by their national research councils. JCAMJ is the recipient of an Australian Research Council Future Fellowship (FT140101082), and also acknowledges support from an Australian Research Council Discovery Grant (FT140101082). EK acknowledges fund from an NWO Vidi fellowship (FT140101082), and also acknowledges support from the Hungarian Scientific Research Fund (OTKA NN110333). SC acknowledges financial support from the UnivEarthS Labex program of Sorbonne Paris Cité (ANR-10-LABX-0023 and ANR-11-IDEX-0005-02).

REFERENCES

Arai A., Uemura M., Sasada M., Trushkin S. A., Ueda Y., Taka-hashi H., Kawabata K. S., Yamanaka M., Nagaee O., Ikejiri Y. e. a., 2009, PASJ, 61, L1 Arnaud K. A., 1996, in G. H. Jacoby & J. Barnes ed., Astronomical Data Analysis Software and Systems V. Vol. 101 of Astronomical Society of the Pacific Conference Series, XSPEC: The First Ten Years. p. 17 Bachetti M., Harrison F. A., Walton D. J., Greisenfette B. W., Chakraborty D., Fürst F., Barret D., Beloborodov A., Boggs S. E., Christensen F. E., Craig W. W., Fabian A. C., et. al. 2014, Nature, 514, 202 Bachetti M., Rana V., Walton D. J., Barret D., Harrison F. A., Boggs S. E., Christensen F. E., Craig W. W., Fabian A. C., Fürst F., Greisenfette B. W., Hailey C. J., Harrison F. A., et. al. 2013, ApJ, 778, 163 Belczynski K., Bulik T., Fryer C. L., 1995, in J. A. Zensus, P. J. Napier ed., Very Long Baseline Interferometry and the VLBA Vol. 82 of Astronomical Society of the Pacific Conference Series, VLBA Data Reduction in Practice. p. 227 Dickey J. M., Lockman F. J., 1990, ARA&A, 28, 215 Egorov O. V., Lozinskaya T. A., Moiseev A. V., 2013, MNRAS, 429, 1450 Farrell S. A., Webb N. A., Barret D., Godet O., Rodrigues J. M., 2009, Nature, 460, 73 Fender R. P., Homan J., Belloni T. M., 2009, MNRAS, 396, 1370 Feng H., Kaaret P., 2010, ApJL, 712, L169 Fryer C. L., 1999, ApJ, 522, 413 Gan Z., Yuan F., Ostriker J. P., Ciotti L., Novak G. S., 2014, ApJ, 789, 150 Gladstone J. C., Roberts T. P., Done C., 2009, MNRAS, 397, 1836 Goas M. R., Roberts T. P., Reeves J. N., Uttley P., 2006, MNRAS, 365, 191 Greisen E. W., 2003, Information Handling in Astronomy - Historical Vistas, 285, 109 Grisé F., Kaaret P., Corbel S., Feng H., Cseh D., Tao L., 2012, ApJ, 745, 123 Grisé F., Kaaret P., Feng H., Kajava J. J. E., Farrell S. A., 2010, ApJL, 724, L148 Heinz S., 2006, ApJ, 636, 316 Jonker P. G., Torres M. A. P., Fabian A. C., Heida M., Minniti G., Pooley D., 2010, MNRAS, 407, 645 Kaaret P., Corbel S., 2009, ApJ, 697, 950 Kaaret P., Corbel S., Prestwich A. H., Zezas A., 2003, Science, 299, 365 Kaaret P., Ward M. J., Zezas A., 2004, MNRAS, 351, L83 Kajava J. J. E., Poutanen J., Farrell S. A., Grisé F., Kaaret P., 2012, MNRAS, 422, 990 Karachentsev I. D., Dolphin A. E., Geisler D., Grebel E. K., Guhathakurta P., Hodge P. W., Karachentseva V. E., Sarajedini A., Seitzer P., Sharina M. E., 2002, A&A, 383, 125 King A., Pounds K., 2015, 2015arXiv150305206K Kovalev Y. Y., Kellermann K. I., Lister M. L., Homan D. C., Vermeulen R. C., Cohen M. H., Ros E., Kardel M., Lobanov A. P., Zensus J. A., Kardashev N. S., Gurvits L. I., Aller M. F., Aller H. D., 2005, AJ, 130, 2473 Lehmann I., Becker T., Fabrika S., Roth M., Miyaji T., Afanasiev V., Sholukhova O., Sánchez S. F., Greiner J., Hasinger G., Costantini E., Surkov A., Burenkov A., 2005, Astronomy, 431, 847 Liu J.-F., Bregman J. N., Bai Y., Justham S., Crowther P., 2013, Nature, 503, 500 Longair M. S., 1994, High energy astrophysics. Vol.2: Stars, the galaxy and the interstellar medium. Cambridge: Cambridge University Press, —c1994, 2nd ed. Mapelli M., Bressan A., 2013, MNRAS, 430, 3120 Mapelli M., Colpi M., Zampieri L., 2009, MNRAS, 395, L71 Marlowe H., Kaaret P., Lang C., Feng H., Grisé F., Miller J., Cseh D., Corbel S., 2014, MNRAS, 439, L1 Diamond P. J., 1995, in J. A. Zensus, P. J. Napier ed., Very Long Baseline Interferometry and the VLBA Vol. 82 of Astronomical Society of the Pacific Conference Series, VLBA Data Reduction in Practice. p. 227
A jet ejection of the ULX Holmberg II X-1

Middleton M. J., Miller-Jones J. C. A., Markoff S., Fender R., Henze M., Hurley-Walker N., Scaife A. M. M., Roberts T. P., Walton D., Carpenter J., Macquart J.-P. e. a., 2013, Nature, 493, 187

Middleton M. J., Walton D. J., Roberts T. P., Heil L., 2014, MNRAS, 438, L51

Miller N. A., Mushotzky R. F., Neff S. G., 2005, ApJL, 623, L109

Motch C., Pakull M. W., Soria R., Grišé F., Pietrzyński G., 2014, Nature, 514, 198

Pacucci F., Ferrara A., 2015, ArXiv e-prints

Pakull M. W., Grišé F., 2008, in R. M. Bandyopadhyay, S. Wachter, D. Gelino, & C. R. Gelino ed., A Population Explosion: The Nature & Evolution of X-ray Binaries in Diverse Environments Vol. 1010 of American Institute of Physics Conference Series, Ultraluminous X-ray Sources: Beambags and Optical Counterparts. pp 303–307

Pakull M. W., Mirioni L., 2002, ArXiv Astrophysics e-prints, 0202488

Pakull M. W., Soria R., Motch C., 2010, Nature, 466, 209

Pasham D. R., Strohmayer T. E., Mushotzky R. F., 2014, Nat, 513, 74

Ponti G., Fender R. P., Begelman M. C., Dunn R. J. H., Neilsen J., Coriat M., 2012, MNRAS, p. L417

Poutanen J., Fabrika S., Valeev A. F., Sholukhova O., Greiner J., 2013, MNRAS, 432, 506

Poutanen J., Lipunova G., Fabrika S., Butkevich A. G., Abolmasov P., 2007, MNRAS, 377, 1187

Prestwich A. H., Tsantaki M., Zezas A., Jackson F., Roberts T. P., Foltz R., Linden T., Kalogera V., 2013, ApJ, 769, 92

Rana V., Harrison F. A., Bachetti M., Walton D. J., Furst F., Barret D., Miller J. M., Fabian A. C., Boggs S. E., Christensen F. C., Craig W. W., Grefenstette B. W., Hailey C. J., Madsen K. K., Ptak A. F., Stern D., Webb N. A., Zhang W. W., 2015, ApJ, 799, 121

Readhead A. C. S., 1994, ApJ, 426, 51

Ruston A., Spencer R., Fender R., Pooley G., 2010, A&A, 524, A29

Shepherd M. C., Pearson T. J., Taylor G. B., 1994, in Bulletin of the American Astronomical Society Vol. 26 of Bulletin of the American Astronomical Society, DIFMAP: an interactive program for synthesis imaging.. pp 987–989

Soria R., Long K. S., Blair W. P., Godfrey L., Kuntz K. D., Lenc E., Stockdale C., Winkler P. F., 2014, Science, 343, 1330

Stobbart A.-M., Roberts T. P., Wilms J., 2006, MNRAS, 368, 397

Tombesi F., Melendez M., Veilleux S., Reeves J. N., Gonzalez-Alfonso E., Reynolds C. S., 2015, 2015arXiv150107664T

van Dyk S. D., Sramek R. A., Weiler K. W., Hyman S. D., Virden R. E., 1994, ApJL, 425, L77

Volonteri M., Silk J., 2014, 2014arXiv1401.3513V

Walton D. J., Harrison F. A., Greffenstette B. W., Miller J. M., Bachetti M., Barret D., Boggs S. E., Christensen F. E., Craig W. W., Fabian A. C., Fuerst F., Hailey C. J., Madsen K. K., Parker M. L., Ptak A., Rana V., Stern D., Webb N., Zhang W. W., 2014, ApJ, 793, 21

Walton D. J., Middleton M. J., Rana V., Miller J. M., Harrison F. A., Fabian A. C., Bachetti M., Barret D., Boggs S. E., Christensen F. E., Craig W. W., Fuerst F., Greffenstette B. W., Hailey C. J., Madsen K. K., Stern D., Zhang W., 2015, ArXiv e-prints

Walton D. J., Miller J. M., Harrison F. A., Fabian A. C., Roberts T. P., Middleton M. J., Reis R. C., 2013, ApJL, 773, L9

Walton D. J., Miller J. M., Reis R. C., Fabian A. C., 2012, MNRAS, 426, 473

Webb N., Cseh D., Lenc E., Godet O., Barret D., Corbel S., Farrell S., Fender R., Gehrels N., Heywood I., 2012, Science, 337, 554

Zampieri L., Roberts T. P., 2009, MNRAS, 400, 677