A giant ring-like structure at 0.78 < z < 0.86 displayed by GRBs

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ABSTRACT

According to the cosmological principle (CP), Universal large-scale structure is homogeneous and isotropic. The observable Universe, however, shows complex structures even on very large scales. The recent discoveries of structures significantly exceeding the transition scale of 370 Mpc pose a challenge to the CP. We report here the discovery of the largest regular formation in the observable Universe; a ring with a diameter of 1720 Mpc, displayed by 9 gamma-ray bursts (GRBs), exceeding by a factor of 5 the transition scale to the homogeneous and isotropic distribution. The ring has a major diameter of 43° and a minor diameter of 30° at a distance of 2770 Mpc in the 0.78 < z < 0.86 redshift range, with a probability of 2 × 10−6 of being the result of a random fluctuation in the GRB count rate. Evidence suggests that this feature is the projection of a shell on to the plane of the sky. Voids and string-like formations are common outcomes of large-scale structure. However, these structures have maximum sizes of 150 Mpc, which are an order of magnitude smaller than the observed GRB ring diameter. Evidence in support of the shell interpretation requires that temporal information of the transient GRBs be included in the analysis. This ring-shaped feature is large enough to contradict the CP. The physical mechanism responsible for causing it is unknown.

Key words: gamma-ray burst: general – cosmology: observations – large-scale structure of Universe.

1 INTRODUCTION

Quasars are well suited for mapping out the large-scale distribution of matter in the Universe, due to their very high luminosities and preferentially large redshifts. Quasars are associated by groups and poor clusters of galaxies (Heimamäki et al. 2005; Lietzen et al. 2009) and can be observed even when the underlying galaxies are faint and difficult to detect. When quasars cluster, they identify considerable amounts of underlying matter, such that quasar clusters have been used to detect matter clustered on very large scales. Some of this matter is clustered on scales equal to or exceeding that of the Sloan Great Wall (Gott et al. 2005).

A number of large quasar groups (LQGs) have been identified in recent years; each one mapping out large amounts of much fainter matter. After Webster (1982) found a group of four quasars at z = 0.37 with a size of about 100 Mpc, having a low probability of being a chance alignment, Komberg, Kravtsov & Lukash (1994) identified strong clustering in the quasar distribution at scales less than 20h−1 Mpc, and defined LQGs using a well-known cluster analysis technique. Subsequently, Komberg, Kravtsov & Lukash (1996) identified additional LQGs, and Komberg & Lukash (1998) reported a new finding of eleven LQGs based on systematic cluster analysis. The sizes of these clusters ranged from 70 to 160 h−1 Mpc. Newman et al. (1998a,b) later discovered a 150 h−1 Mpc group of 13 quasars at median redshift z ≃ 1.51. Williger et al. (2002) mapped 18 quasars spanning ≃ 5° × 2.5° on the sky, with a quasar spatial overdensity 6–10 times greater than the mean. Haberzettl et al. (2009) investigated two sheet-like structures of galaxies at z = 0.8 and 1.3 spanning 150 h−1 comoving Mpc embedded in LQGs extending over at least 200 h−1 Mpc. Haines, Campusano & Clowes (2004) reported the finding of two large-scale structures of galaxies in a 40 × 35 arcmin2 field embedded in a 25 deg2 area containing two 100 Mpc scale structures of quasars. As the identified scales of quasar clusters became larger, Clowes et al. (2012) found two relatively close LQGs at z ≃ 1.2. The characteristic sizes of these two LQGs, ≃350–400 Mpc, and appear to be only marginally consistent with the scale of homogeneity in the concordance cosmology. Recently, Clowes et al. (2013) uncovered the Huge-LQG with a long dimension of ≃1240 Mpc (1240 × 640 × 370 Mpc). Until recently, this was the largest known structure in the Universe. Using a friend of friend algorithm Einasto et al. (2014) found that the linking length l = 70 h−1 Mpc three systems from their quasar
Horváth, Hakkila & Bagoly (2013, 2014) announced the discovery of a larger Universal structure than the Huge-LQG by analysing the spatial distribution of gamma-ray bursts (GRBs). The 3000 Mpc size of this structure exceeds the size of the Sloan Great Wall by a factor of about 6; this is currently the largest known universal structure.

Unlike quasars, GRBs are short-lived cosmic transients spanning milliseconds to hundreds of seconds (see the review paper by Mészáros 2006). Due to their immense luminosities, GRBs can be observed at very large cosmological distances. Their hosts are typically metal-poor galaxies of intermediate mass (Savaglio, Glazebrook & Le Borgne 2009; Castro Cerón et al. 2010; Levesque et al. 2010), rather than the massive galaxies in which quasars are generally found. Both quasars and GRBs should map the underlying distribution of universal matter, although the details of their spatial distributions are not necessarily identical. Although the number of known GRBs for which distances have been accurately measured is significantly fewer than the number of known quasars, the surveying techniques used to identify these objects are more homogeneous and better suited for studying structures of large angular size than are quasars.

The existence of an object with a size of several gigaparsecs introduces questions concerning the homogenous and isotropic nature of cosmological models. The great importance of this question necessitates further independent study into the use of GRBs for mapping large-scale universal structures. Our analysis considers the space distribution of a GRB sample having known redshifts for the presence of large-scale anisotropies. The survey we use for this study is available at http://www.astro.caltech.edu/grbox/grbox.php. As of 2013 October, the redshifts of 361 GRBs have been determined.

2 DISTRIBUTION OF GRBS IN \{r, θ, ϕ\} SPACE

According to the cosmological principle (CP), Universal large-scale structure is homogenous and isotropic (Ellis 1975). The WMAP and Planck experiments have revealed that the big bang had these properties in its early expansion phase. The observable Universe, however, shows complex structures even on very large scales. The problem is to find a limiting scale at which the CP is valid.

A number of well-known studies have attempted to find the largest scale on which CP is valid. According to Einasto & Gramm (1993), the available data suggested values \( r(t) = 130 \pm 10 \, h^{-1} \). Yahata et al. (2005) reported on the first result from the clustering analysis of Sloan Digital Sky Survey (SDSS) quasars: the bump in the power spectrum due to the baryon density was not clearly visible, and they concluded that the galaxy distribution was homogeneous on scales larger than 60–70 \( h^{-1} \) Mpc. Tegmark et al. (2006) using luminous red galaxies (LRGs) in the SDSS improved the evidence for spatial flatness (\( \Omega_{\text{tot}} = 1 \)). Liivamägi, Tempel & Saar (2012) have constructed a set of supercluster catalogues for the galaxies from the SDSS survey main and LRG flux-limited samples. Bagla, Yadav & Seshadri (2008) showed that in the concordance model, the fractal dimension makes a rapid transition to values close to 3 at scales between 40 and 100 Mpc. Sarkar et al. (2009) found the galaxy distribution to be homogeneous at length-scales greater than 70 \( h^{-1} \) Mpc, and Yadav, Bagla & Khandai (2010) estimated the upper limit to the scale of homogeneity as being close to 260 \( h^{-1} \) Mpc for the\( \Lambda \) cold dark matter (\( \Lambda \text{CDM} \)) model. Söchting et al. (2012) studied the Ultra Deep Catalogue of Galaxy structures; the cluster catalogue contains 1780 structures covering the redshift range 0.2 < \( z < 3.0 \), spanning 3 orders of magnitude in luminosity (\( 10^8 < L_\odot < 5 \times 10^{11} L_\odot \)) and richness from eight to hundreds of galaxies. These results supported the validity of CP.

Assuming the validity of CP, a homogeneous isotropic model and a standard \( \Lambda \text{CDM} \) cosmology (\( \Omega_k = 0.7, \Omega_M = 0.3, h = 0.7 \), representing an Euclidean space with \( \Omega_k + \Omega_M = 0.7 + 0.3 = 1 \) the line element in the 4D space–time is given by

\[
dl^2 = R(t)^2 (dr^2 + r^2 d\theta^2 + r^2 \sin \theta^2 d\phi^2) - c^2 dt^2. \tag{1}
\]

The variables in the equation have their conventional meaning. The change of the spatial part of \( dl^2 \) line element in course of time is given by

\[
ds^2 = dr^2 + r^2 d\theta^2 + r^2 \sin \theta^2 d\phi^2 \tag{2}
\]

is independent of the time. Any event in the 4D space–time has a footprint in the \( \{r, \theta, \phi\} \) space, where \( r \) can be computed from the observed redshift \( z \) and the angular coordinates are given by the observations. In astronomy, the angular coordinates are usually concretized in equatorial or Galactic systems. The \( r \) distance is measured by the comoving distance defined in the Euclidean case by

\[
r(z) = \frac{c}{H_0} \int_0^z \frac{dz'}{\sqrt{\Omega_M (1 + z'^2) + \Omega_\Lambda}}, \tag{3}
\]

where \( c \) is the speed of light and \( H_0 \) is the Hubble constant.

The distribution of GRBs in the \( \{r, \theta, \phi\} \) coordinate system can be constructed by assuming some universal formation history, along with spatial homogeneity and isotropy. This theoretical distribution, however, cannot be observed directly because the observations are biased by selection effects. There are several factors influencing the probability that a GRB is detected. The limit of the instrumental sensitivity is such a factor, as GRBs below this threshold cannot be detected. The probability of detecting a GRB depends also on the observational strategy of the satellite.

The GRBs in our sample having known redshift were detected by different satellites using different observational strategies. The method of observation results in different detection probabilities (known as the exposure function) since each satellite spends different time durations observing various parts of the sky. In principle, this exposure function can be reconstructed from the log of observations made by each satellite. The GRB redshift can be obtained either from spectral observation of the GRB afterglow or of the host galaxy if it can be localized. The exposure is a function of GRB brightness and not of GRB redshift.

Redshift measurements have their own biases. These depend on the optical brightness of the afterglow or the host galaxy, depending on which is observed. The most significant factor influencing the possibility of observation is extinction caused by the Galactic foreground. This bias seriously influences the probability of redshift measurement. However, galactic extinction has a known distribution and can be estimated at cosmological distances for any angular position on the sky.

Most GRB afterglows and host galaxies are optically faint so one needs large-aperture telescopes in order to measure GRB redshifts. Northern and Southern hemisphere telescopes used to make these measurements are located at mid-latitudes; consequently the chance of getting the necessary observing time is higher in the winter than in the summer since the night is much longer during this season. That part of the sky which is accessible in the winter season has a higher probability for determining a GRB redshift. This bias is equalized when observations are made over many observing
seasons. However, since host galaxy measurements can be made at any time after the GRB is observed, this effect is less important when measuring the redshift from the host.

The net effect of all these factors can be expressed by the following formula:

\[ f_{\text{obs}}(\theta, \varphi, z) = T_{\text{int}}(\theta, \varphi) E_{\text{ex}}(\theta, \varphi) E_{\text{xp}}(\theta, \varphi) f_{\text{int}}(z). \]  \hspace{1cm} (4)

The \( T_{\text{int}} \) telescope time and \( E_{\text{xp}} \) exposure function factors do not depend on the redshift. As to the \( E_{\text{ex}} \) extinction, however, one may have some concerns. If the observable optical brightness depends on the distance, which seems to be a quite reasonable assumption, then the measured \( z \) values close to the Galactic equator could be systematically lower. Our sample of 361 GRBs, however, does not show this effect. One may assume, therefore, that the observed distribution of GRBs in the \( \{ r, \theta, \varphi \} \) space can be written in the form of

\[ f_{\text{obs}}(\theta, \varphi, z) = g(\theta, \varphi) f_{\text{int}}(z). \]  \hspace{1cm} (5)

In deriving this equality, it is assumed that \( f_{\text{int}} \) does not depend on the \( \theta, \varphi \) angular coordinates, due to isotropy.

### 3 TESTING THE ANGULAR ISOTROPY OF THE GRB DISTRIBUTION

A number of approaches have been developed for studying the non-random departure from the homogeneous isotropic distribution of matter in the universe. Each method is sensitive to different forms of the departure of the cosmic matter density from the homogeneous isotropic case. These tests have been previously applied to galaxy and quasar distributions. Clowes, Iovino & Shaver (1987) described a three-dimensional clustering analysis of 1100 ‘high-probability’ quasar candidates occupying the assigned-redshift band of 1.8–2.4. Icke & van de Weygaert (1991) presented a geometrical model, making use the Voronoi foam, for the asymptotic distribution of the cosmic mass on 10–200 Mpc scales. Graham, Clowes & Campusano (1995) applied a graph theoretical method, using the minimal spanning tree (MST), to find candidates for quasar superstructures in quasar surveys. Doroshkevich et al. (2004) used the MST technique to extract sets of filaments, wall-like structures, galaxy groups and rich clusters. Platen et al. (2011) studied the linear Delaunay Tessellation Field Estimator (DTFE), its higher order equivalent Natural Neighbour Field Estimator (NNFE) and a version of the Krigeing interpolation. The DTFE, NNFE and Krigeing approaches had largely similar density and topology error behaviours. Zhang, Springel & Yang (2010) introduced a method for constructing galaxy density fields based on DTFE. Recently, Kitaura (2013, 2014) presented the KIGEN-approach which allows for the first time for self-consistent phase-space reconstructions from galaxy redshift data.

Even if CP is correct, the distribution of GRBs in the comoving frame will likely be heterogeneous due to the cosmic history of structure formation. However, the isotropy of the angular distribution can remain unchanged. Based on the data obtained by the CGRO Burst and Transient Source Experiment (BATSE) experiment, Briggs (1993) concluded that angular distribution of GRBs is isotropic on a large scale. In contrast, later studies performed on larger samples suggest that this is true only for bursts of long duration (>10 s) rather than subclasses of GRBs having different durations and presumably different progenitors (Balazs, Meszaros & Horvath 1998; Balazs et al. 1999; Meszaros et al. 2000; Litvin et al. 2001; Tanvir et al. 2005; Vavrek et al. 2008).

### 3.1 Testing the isotropy based on conditional probabilities

The validity of equation (5) can be tested by some suitably chosen statistical procedure. The equation indicates that the distribution of the angular coordinates is independent of the \( z \) redshift. By definition, this independence also means that the \( g(\theta, \varphi | z) \) conditional probability density of the angular distribution, assuming that \( z \) is given, does not depend on \( z \), i.e. the joint probability density of the angular and redshift variables can be written in the form

\[ f_{\text{obs}}(\theta, \varphi, z) = g(\theta, \varphi) f_{\text{int}}(z). \]  \hspace{1cm} (6)

Let us consider an \( \{ \theta_i, \varphi_i, z_i ; i = 1, 2, \ldots, n \} \) observed sample of GRB positional and redshift data. The validity of equation (6) can be tested on this data set in different ways. One approach is to test whether or not the conditional probability is independent of the condition. This is the approach used by Horvath et al., who split the sample into \( k \) subsamples according to the \( z \)-characteristics and tested whether or not the subsamples could originate from the same \( g(\theta, \varphi) \). This is the null hypothesis to be tested. If the null hypothesis is true, then the way in which the sample is subdivided by \( z \) is not crucial. A reasonable approach is to select the subsamples so that they have equal numbers of GRBs; this ensures the same statistical properties of the subsamples. The number of subsamples \( k \) is somewhat arbitrary.

Horvath et al. selected \( k \) different radial bins and performed several tests for sample isotropy on each bin. The tests singled out the slice at \( 1.6 < z < 2.1 \) as having a statistically significant sample anisotropy, which was due to a large cluster of GRBs in the northern Galactic sky.

### 3.2 Testing the isotropy of the GRB distribution using joint probability factoring

It is important to independently test the significance of the Horvath et al. result using alternative statistical methods. We choose to regard the GRB distribution in terms of equation (6). The right-hand side of equation (6) describes a factorization of the \( f_{\text{obs}}(\theta, \varphi, z) \) joint probability density in terms of the angular and redshift variables. To test the validity of this factorization we proceed in the following way: we consider again the sample of \( \{ \theta_i, \varphi_i, z_i ; i = 1, 2, \ldots, n \} \) mentioned above. If the factorization is valid (e.g. if the angular distribution is independent of redshift), then the sample is invariant under a random resampling of the \( z \) variable while keeping the angular coordinates unchanged. Using this approach, we get a new sample that is statistically equivalent to the original one, assuming that the joint probability density factorization is valid.

The sample coming from the \( f_{\text{obs}}(\theta, \varphi, z) \) joint probability density is three-dimensional, so the task at hand is a comparison of three-dimensional samples. One way in which the sample can deviate from isotropy is through the presence of one or more density enhancements and/or decrements. For this type of density perturbation \( f_{\text{int}} \) is also dependent on the angular coordinates and the factorization is no longer valid. In this case, the resampling of the redshifts may change \( f_{\text{obs}} \), as well. Comparing the original observed distribution with the resampled one can identify the presence of density perturbations with respect to the isotropic case.

To calculate the Euclidean distances in a simple way, we introduce \( x_c, y_c, z_c \). Descartes coordinates in the comoving \( r, \theta, \varphi \) frame. Using Galactic coordinates, one obtains

\[ x_c = r \cos(b) \cos(l) \]  \hspace{1cm} (7)
Depending on the standardized logarithmic density on the $\sigma / 2$ follows $\sigma \chi$ lines at about 2800 Mpc. These dots may represent constant and 2 $k$ lines in the 4000–6000 Mpc.

The result is displayed in Fig. 1. The dependence of the density variance on the 4th order of the nearest neighbour. Note the rapid decrease of variance at lower $k$ and the much shallower decrease for larger values.

$$y_k = r \cos(b) \sin(l)$$

$$z_k = r \sin(b) \quad (8)$$

Having a sample of size $N$ the number of objects in a differential volume $dV$ can be written using Descartes coordinates in the form $dN = v(x, y, z)dx dy dz$, where $v$ denotes the spatial density of the objects. It is clear from this formula that $v(x, y, z) = N_{\text{cell}}(x, y, z)$. Obviously, $v(x, y, z) = dV/dV$.

A trivial estimate of $v$ is obtained by counting $dN$ in a given $dV$. Fixing $dV$, the variance of $v$ is given by $dV$. An alternative approach is to keep $dV$ constant and look for the appropriate $dV$. This approach can be realized by computing the distance to the $k$th nearest neighbour in the sample. The distance is the conventional Euclidean one in our case. Proceeding in this way, $v(x, y, z) = 3(k + 1)/(4\pi r_k^3)$, where $r_k$ is the distance to the $k$th nearest neighbour. In the following subsection, we adopt this approach for the GRB sample.

### 3.3 $k$th nearest neighbour statistics for the GRB data

In order to find the $k$th nearest neighbours, we use the `knn.dist(x, k)` procedure in the `FNN` library of the R statistical package. Given $k$, the procedure computes the distances of the $x$ sample elements up to the $k$th order. Resampling the data $r$ can proceed by using the `sample(x, n)` procedure of $r$ where $x$ refers to the data set being resampled and $n$ is the size of the new sample (which is identical to the original one in our case). We have made 10 000 resamplings of the comoving distance sample and computed the densities choosing $k = 1, 2, \ldots, 20$ for the nearest neighbour distances. The densities obtained from the resampled data enable us to compute the mean density and its variance at each of the sample distances.

At this point, the value of $k$ is quite arbitrary. Selecting a small value of $k$ might make this procedure sensitive to small-scale disturbances. However, the variance of the estimated density on this scale is large. In contrast, the variance of the density is lower for larger values of $k$ but the density fluctuations of smaller scale might be smeared out. In order to make a reasonable compromise between $k$ and the variance of the estimated density, we compute the mean variance in the function of $k$. The result is displayed in Fig. 1.

The figure helps us to estimate the optimal selection of the $k$ value. Up to $k = 5$, the variance drops rapidly and starting about $k = 8$ its change is much shallower. We select the values of $k = 8, 10, 12, 14$ and the variance of the density at every GRB location in the sample. Using the variance and the mean of the density, we calculate the standardized zero mean and unit variance values of the density for all points of the sample. In Fig. 2, we display a scatter plot between the standardized logarithmic density and the comoving distance for the GRB data.

The null hypothesis in this case is the assumption that all spatial density enhancements of GRBs are produced purely by random fluctuations. Assuming the validity of the null hypothesis one has to compute the probability of the density enhancement in question. We perform a series of Kolmogorov–Smirnow (KS) tests and confirm that the standardized logarithmic densities displayed in Fig. 2 follow a Gaussian distribution. Denoting the logarithmic density by $\varphi$, the probability that the $i$th element’s value is a density enhancement is $p_i \propto \exp(-\varphi_i^2/2)$ and the log-likelihood function has the form

$$L = -\frac{1}{2} \sum_{i=1}^{n} \varphi_i^2 + \text{const.} \quad (10)$$

The summation in equation (10) yields a $x^2_\nu$ variable with $n$ degrees of freedom. All density fluctuations are restricted within certain

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1. [http://cran.r-project.org](http://cran.r-project.org)
distance ranges. To get a likelihood function sensitive to a given density enhancement and/or deficit, we have to sum up successive points in equation (10). Taking and summing up $k$ successive points within a distance range we get a $\chi^2$ variable having $k$ degrees of freedom. All density values are calculated from a fixed number of nearest neighbours in the sample. We choose for $k$ the order of the nearest value from which the density is calculated.

There is some arbitrariness to this procedure. The $k$ orders of the nearest neighbours have been selected somewhat by insight. Nevertheless, the calculated densities correlate strongly. This property enables us to concentrate the densities into one variable.

The correlation matrix in Table 1 reveals a strong correlation between nearest neighbour estimation of densities of the orders $k = 8$, $k = 10$, $k = 12$ and $k = 14$. To get a joint variable, we perform principal component analysis (PCA) on the standardized logarithmic densities used above. To get the PCs, we use the princomp() procedure from R’s stats library.

By running this procedure, we obtain the eigenvalues given in Table 2. The eigenvalues indicate the variance of PC’s obtained. We infer from the variances that the first PC describes 91 per cent of the total variance. We assume, therefore, that the information from the spatial density is concentrated into this variable. From the first PC, we compute the $\chi^2$ variables for df = 8, df = 10, df = 12 and df = 14 degrees of freedom. The results are displayed in Fig. 3.

### Table 1. Correlation between logarithmic densities computed from $k$th orders of nearest neighbours.

| Order | $k$ = 8 | $k$ = 10 | $k$ = 12 | $k$ = 14 |
|-------|---------|---------|---------|---------|
| $k$ = 8 | 1.000 | 0.914 | 0.540 | 0.706 |
| $k$ = 10 | 0.914 | 1.000 | 0.515 | 0.661 |
| $k$ = 12 | 0.540 | 0.515 | 1.000 | 0.743 |
| $k$ = 14 | 0.706 | 0.661 | 0.743 | 1.000 |

### Table 2. Eigenvalues resulting from the PCA. The values demonstrate clearly that only the first PC (marked in bold face) is significant.

| PC1 | PC2 | PC3 | PC4 |
|-----|-----|-----|-----|
| Eigen val. | 3.517 | 0.222 | 0.077 | 0.049 |
| St. dev. | 1.875 | 0.472 | 0.279 | 0.221 |

4 DISCOVERY AND NATURE OF THE GRB RING

In all frames in Fig. 3, there is a strong peak exceeding the 99.9 per cent significance level (99.95, 99.93, 99.96, 99.97 per cent at $k = 8$, $k = 10$, $k = 12$ and $k = 14$, respectively) at about 2800 Mpc corresponding to a group of outlying points in Fig. 2. This appears to indicate the presence of some real density enhancement.

#### 4.1 Discovery of the ring

We assume that GRBs making some contribution to the highly significant peak shown in Fig. 3 lie within the full width at half-maximum (FWHM) angular distance of the peak. Their angular distribution is shown in Fig. 4. The most conspicuous feature in all of the frames is a ring-like structure in the lower left side of the frames. The redshift, distance and Galactic coordinates of the GRBs displaying the ring are given in Table 3. Using the data listed in the table, one can calculate the mean redshift and distance of the ring, along with the standard deviations of these variables, yielding $z = 0.822, \sigma_z = 0.025$ and $d_c = 2770$ Mpc; $\sigma_d = 65$ Mpc.

By definition, the true characteristic physical size $D$ of the object can be obtained from the $D = \Theta \times d_c = \Theta \times d_c/(1 + z)$ relation, where $\Theta$ is the mean angular size and $d_c$ is the angular...
distance. Substituting the corresponding values obtained above one gets \( D = 944 \) Mpc, corresponding to 1720 Mpc in the comoving frame.

### 4.2 Verification of the ring structure

In Section 4.1, we claim to find a regular structure in the shape of a ring. However, the form of this structure is thus far based only on a visual impression. In the following discussion, we try to give a quantitative value supporting the sensibility of this subjective impression.

The procedure we used to obtain the very low probability of this density enhancement only by chance is not sensitive to the true shape of this clustering. Assuming this clustering to be real, we may compute the probability of getting a ring-like structure only by chance by assuming some concrete space distribution for the objects. We make these calculations for the cases of (a) a homogeneous sphere and (b) a shell (for more details see the Appendix).

The probability of observing a ring-like structure is \( p = 0.2 \) in the case of a shell but it is only \( p = 4 \times 10^{-3} \) for a homogeneous sphere. It is worth noting that real space distributions of cluster members generally concentrate more strongly towards the centre than do the elements of a homogeneous sphere. Therefore, this probability can be taken as an upper bound for a probability of obtaining a ring shape purely by chance.

Combining this latter probability with that of observing the clustering purely by chance, we obtain a value of \( p = 2 \times 10^{-6} \) for observing a ring entirely by chance. Thus, despite the large angular size of the extended GRB cluster, we find evidence that the cluster represents a large extended ring in the 0.78 \( \leq z \leq 0.86 \) redshift range.

### 4.3 The physical nature of the ring

If we assume that the ring represents a real structure, then we can speculate about its nature and origins. Perhaps a simple explanation is that it indicates the presence of a ring-like cosmic string. This would indicate that it is a large-scale component of the cosmic web, representing the characteristic spatial distribution of the objects in the Universe. The main difficulty with this simple explanation lies with the uniformity of the redshifts (distances) along the object, indicating that we must be seeing the ring nearly face-on. This possibility cannot be excluded, but alternative explanations are also worth of considering.

GRBs are short-lived transient phenomena. The GRBs that compose the ring, along with their redshifts, were collected over a period of about 10 yr. The number of observed events is determined by the time frequency of such events in a given host galaxy. However, the total number of hosts in the region containing the GRB ring and not having burst events during the observation period must be much greater relative to those which were observed.

The number of observed GRBs should be proportional to the number of progenitors in the same region, although the spatial stellar mass density is not necessarily proportional to the spatial number density of progenitors. Namely, the progenitors for the majority of GRBs are thought to be short-lived 20–40 M\(_{\odot}\) stars, and as such their presence should be strongly dependent on the star formation activity in their host galaxies. Thus, our knowledge of the underlying mass distribution is sensitive to assumptions about star formation within the ring galaxies.

#### 4.3.1 Mass of the ring

In order to estimate the mass of the ring structure, we make two extreme assumptions providing lower and upper bounds for the mass. We get a lower bound for the mass by assuming that the general spatial stellar mass density is the same in the field and in the ring’s region and only the star formation and consequently the GRB formation rate is higher here. We get an upper bound for the mass by supposing a strict proportionality between the stellar mass density and the number density of the progenitors. For both estimates, we need to know the local stellar density.

Several recent studies have attempted to determine the stellar (baryonic) mass density and its relation to the total Universal mass density. Bahcall & Kullier (2014) determined the stellar mass fraction and found it to be nearly constant on all scales above 300 h\(^{-1}\) kpc, with \( M^*/M_{\text{tot}} \approx 0.01 \pm 0.004 \). Le Fèvre et al. (2013, 2014) issued the VIMOS VLT Deep Survey (VVDS) final and public data release offering an excellent opportunity to revisit galaxy evolution. The VVDS is a comprehensive survey of the distant universe, covering all epochs since \( z \sim 6 \). From this, Davidzon et al. (2013) measured the evolution of the galaxy stellar mass function from \( z = 1.3 \) to \( z = 0.5 \) using the first 53 608 redshifts.

Marulli et al. (2013) investigated the dependence of galaxy clustering on luminosity and stellar mass in the redshift range \( 0.5 < z < 1.1 \) using the ongoing VIMOS Public Extragalactic Survey. Based on their sample of 10 095 galaxies, Driver et al. (2007) estimated the stellar mass densities at redshift zero amounting to \( 8.6 h^{-1} \) \( 0.6 \times 10^8 M_{\odot} h^{-1} \text{Mpc}^{-3} \). We use this local value as our measure of the mass density in the comoving frame and use it for our subsequent calculations.

We assign a volume to the ring by computing the convex hull (CH) of the points representing the GRBs in the rest frame. Using the QHULL\(^2\) program, we obtain a value of \( 1.9 \times 10^8 \) Mpc\(^3\) for the volume of the CH. Supposing that the stellar mass density is the same in the ring’s region (i.e. within the CH) as in the field and only the number of progenitors is enhanced, we compute a mass of \( 2.3 \times 10^{17} M_{\odot} \) inside the CH.

Alternatively, if we assume that the fraction of progenitors is the same along the ring as it is in the field, then the total stellar mass in the volume of a shell with 2770 Mpc radius and 200 Mpc thickness (the observed distance range of the GRBs in the ring) is \( 2.2 \times 10^{19} M_{\odot} \). Since the number of GRBs making up the ring is about half the total observed in the shell, we get a mass of \( 1.1 \times 10^{19} M_{\odot} \) within the ring’s CH.

Supposing a strict proportionality between the spatial densities of the number of GRB progenitors and stellar masses we estimate a factor of 50 times more mass than that which is obtained above using the field value within the CH. In the case of a homogeneous mass distribution within the CH this implies an overdensity by a factor of 50 compared to the field.

In reality, however, the overdensity within the CH appears to be concentrated in the outer half of its volume in order to produce a ring-like distribution, suggesting an overdensity enhanced by a factor of more than 100. This high value appears to be unrealistic. To resolve this contradiction, the proportion of progenitors in the stellar mass has to be increased by at least an order of magnitude and the stellar mass density has to be decreased by the same factor in the outer half volume of the CH. This results in a value of \( 1 \times 10^{18} M_{\odot} \).

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\(^2\) [http://www.qhull.org/](http://www.qhull.org/) QHULL implements the Quickhull algorithm for computing the convex hull.
which still represents an overdensity of a factor of 10 suggesting the ring mass is in the range $10^{17}$–$10^{18}$ $M_{\odot}$, depending on the fraction of progenitors in the stellar mass distribution.

4.3.2 The case for a spheroidal structure

To overcome the difficulty caused by the low probability of seeing a ring nearly face-on, one may assume the GRBs populate the surface of a spheroid which we see in projection. To demonstrate that the projection of GRBs uniformly populating the surface of the spheroid really can produce a ring in projection, we make MC simulations displayed in Fig. 5. The simulations show that a ring structure can be obtained easily by projecting a spheroidal shell on to a plane. The probability of observing a ring in this way is much larger than that of observing a ring face-on.

Unfortunately, this approach also faces some problems. Assuming the observed ring is a projection on to a plane, one can calculate the standard deviation of distances of the objects to the observer. A simple calculation shows this standard deviation is about 58 per cent of the radius in case of a sphere. Previously we obtained 1720 Mpc for the diameter of the ring resulting in a 860 Mpc radius with a typical result of $\sigma = 499$ Mpc. The endpoints correspond to some lookback time difference between GRBs detected at the same moment by the observer. The lookback time can be calculated from the following equation:

$$t_L(z) = t_0 \int_0^z \frac{dz'}{\sqrt{\Omega_M (1 + z')^3 + \Omega_\Lambda}},$$

where $t_0 = 1/H_0$ is the Hubble time. Calculating the time difference one obtains $\Delta t_L = 1.9 \times 10^8$ yr. Computing the time difference taking the observed 65 Mpc standard deviation instead of 499 Mpc, one gets only $\Delta t_L = 2.5 \times 10^8$ yr. If the GRBs displaying the observed ring really populate a sphere, then the presence of the low 65 Mpc standard deviation reveals a $2.5 \times 10^8$ yr period in the life of the host galaxy when it is very active in producing GRBs. Furthermore, one has to assume that this happens for all hosts simultaneously. This coordinated activity may happen by some external effect which is responsible for the formation of the sphere.

One can make a similar estimate of the sphere’s mass by assuming that the ring represents a real structure and is not simply a projection. The difference in this case is that the ring mass represents only a fraction of the sphere’s mass. There are seven objects in the ring within the 65 Mpc standard deviation. There is a factor of 7 for getting the number of GRBs within 1 $\sigma$ on the sphere. This number represents 68 per cent of the total number of GRBs on the sphere, i.e. one should multiply the mass obtained for the ring by roughly a factor of 10, yielding $10^{13}$–$10^{14}$ $M_{\odot}$ for the sphere.

4.3.3 Formation of the ring

No matter whether we interpret the spatial structure of the ring as a torus or as a projection of a spheroidal shell, the formation of a structure with this large size and mass provides a real challenge to theoretical interpretations. In addition to the size and the mass of the structure, one has to explain why the GRB activity is much higher along the ring than it typically is in the field.

There is general agreement among researchers that following the early phase of the big bang the initial perturbations evolved into a cosmic web consisting of voids surrounded by string-like structures. A filamentary structure surrounding sphere-like voids is a typical result of gravitational collapse (Centrella & Melott 1983; Icke 1984).

The hierarchy of structures in the density field inside voids is reflected by a similar hierarchy of structures in the velocity field (Aragon-Calvo & Szalay 2013). The void phenomenon is due to the action of two processes: the synchronisation of density perturbations of medium and large scales, and the suppression of galaxy formation in low-density regions (Einasto et al. 2011).

It is generally assumed that the maximal size of these structures is 100–150 Mpc (Frisch et al. 1995; Einasto et al. 1997; Suhhonenko et al. 2011; Aragon-Calvo & Szalay 2013). Quite recently Tully et al. (2014) discovered the local supercluster (the ‘Laniakea’) having a diameter of 320 Mpc. This scale is several times smaller than the estimated 1720 Mpc diameter of the GRB ring, although perturbations on larger scales cannot be excluded. However, Doroshkevich & Klypin (1988) have presented arguments that
perturbations on the 200–300 Mpc scale should be excluded. This value is in a clear contradiction to the existence of the GRB ring and other large observed structures. Resolution of this contradiction is still an open issue.

The existence of the ring, either as a torus or the projection of a spheroidal shell, requires a higher spatial frequency of the progenitors along the ring than in the field. A possible interpretation of the higher fraction of progenitors along the ring is that hosts are still in the formation process at $6.7 \times 10^9$ yr after the big bang. This supports the view that large-scale structure can form and evolve slowly from the initial perturbations (Zeldovich, Einasto & Shandarin 1982; Einasto et al. 2006).

Dark matter must be given a dominant role in large-scale structure theories in order to account quantitatively for the formation and evolution of the cosmic web. That is because the observed distributions of galaxies are inconsistent with gravitational clustering theories and with the formation of super clusters in a wholly gaseous medium (Einasto, Joeveer & Saar 1980). Recent extensive numerical studies that include dark matter indicate that its presence accounts for the basic properties of the cosmic web (Springel et al. 2005; Angulo et al. 2012), and these studies have reproduced the cosmic star formation history and the stellar mass function with some success (Vogelsberger et al. 2013). The very large high-resolution cosmological N-body simulation, the Millennium-XXL or MXXL (Angulo et al. 2012), which uses 303 billion particles, modelled the formation of dark matter structures throughout a 4.1 Gpc box in a ΛCDM cold dark matter cosmology. Kim et al. (2011) presented two large cosmological N-body simulations, called Horizon Run 2 (HR2) and Horizon Run 3 (HR3), made using 6000$^3 = 216$ billions and 7210$^3 = 374$ billion particles, spanning a volume of (7.200 h$^{-1}$ Gpc)$^3$ and (10.815 h$^{-1}$ Gpc)$^3$, respectively. Although these sizes of the simulated volumes were large enough to produce very large structures, accounting for local enhancements corresponding to the size of the GRB ring structure is still an open problem. We address this issue in the next subsection.

4.3.4 Spatial distribution of GRBs and large-scale structure of the Universe

In Section 4.3.3, we noted that very large scale cosmological simulations may account for huge disturbances in the dark matter distribution, in particular for LQGs and the object discovered by Horváth et al. (2013). Nevertheless, in Sections 4.3.1 and 4.3.2 we pointed out that the existence of the GRB ring probably cannot be accounted for a simple enhancement of the underlying baryonic and dark matter density. Presumably, to explain the existence of the ring one needs a coordinated enhanced GRB activity in the responsible host galaxies.

According to a widely accepted view the majority of the observed GRBs are resulted in collapsing high-mass (20–40 M$_\odot$) stars. GRBs are very rare transient phenomena, consequently, they observed spatial distribution is a serious under sampling of the space distribution of galaxies in general. Furthermore, the high-mass stars have short lifetimes, consequently GRBs prefer those galaxy hosts having considerable star-forming activity.

Due to their immense intrinsic brightnesses, GRBs can be detected at large cosmological distances. GRB 090423 has $z = 8.2$, the largest spectroscopically measured redshift (Tanvir et al. 2009). Even though GRBs seriously undersample the matter distribution, they are the only observed objects doing so for the Universe as a whole up to the distance corresponding to the largest measured redshift.

![Histogram of dark matter density (z=0.82)](http://galformod.mpa-garching.mpg.de/portal/mxxl.html)

Figure 6. Probability distribution of the dark matter density in the Millennium simulation. Distribution of dark matter density in the simulation (green), for galaxies in general (red) and for star-forming galaxies (blue). The star-forming galaxies prefer a certain range of underlying dark matter density that differs from that of galaxies in general.

Since there is no complete observational information on the spatial distribution of dark and baryonic matter on the same scale as that of GRBs, one has to use the large-scale simulations of the distribution of the cosmic matter for making such comparisons. We used for this purpose, the publicly available MXXL simulation.3

As we mentioned at the end of Section 4.3.3, the 4.1 Gpc size of the simulated volume is large enough to account for structures with characteristic size of the ring. Since GRBs prefer host galaxies with high star formation activity, we calibrate the dark matter density in XXL to the spatial number density of galaxies having large star-forming rate (SFR) assuming that

$$v_s(x, y, z) = c(q_d)\varrho_d(x, y, z),$$

(12)

where $v_s$ represents the spatial number density of star-forming galaxies and $\varrho_d$ the density of the dark matter. We assume that the $c(q_d)$ conversion factor depends only on $\varrho_d$ but not on the spatial coordinates and that it is identical in the XXL and the Millennium simulation.

We determine the $c(q_d)$ conversion factor using the data available in the Millennium simulation. The GRB ring is located in the 0.78 $z < 0.86$ redshift range, so we select the $z = 0.82$ slice of the simulation. The star-forming galaxies have SFR $> 30$ M$_\odot$ yr$^{-1}$ at this redshift (Perley et al. 2015).

As one can infer from Fig. 6, the selected star-forming galaxies prefer a certain dark matter density range: for densities less than and greater this range such galaxies are uncommon in the sample. This range differs from that of galaxies in general. This may indicate that the spatial distribution displayed by the galaxies in general is not necessary identical with that shown by the GRBs.

After determining the $c(q_d)$ conversion factor, we obtain from equation (6) the $v_s$ spatial number density of the star-forming galaxies in the XXL simulation. Based on this spatial distribution, we generate random samples of sizes comparable to that of the observed GRB frequency. For generating the simulated sample, we use the Markov Chain Monte Carlo (MCMC) method implemented in the metrop($\beta$) procedure available in its mcmc library.

An important issue for using this algorithm is to check whether or not the simulated Markov chain has reached its stationary stage. We check it by computing the auto regression function of the simulated sample by the acf($\beta$) function in $\beta$. We also check the MCMC output by conventional MC.

3 http://galformod.mpa-garching.mpg.de/portal/mxxl.html
The known number of GRBs now exceeds a couple of thousand and is steadily increasing with ongoing observations. Unfortunately, only a fraction of these have measured redshifts. Motivated by the number of known GRBs and by their relationship to star-forming galaxies, we make MCMC simulations of the \( v_z(x, y, z) \) spatial number density of these galaxies from equation (6), getting sample sizes of 1000, 5000, 10 000 and 20 000.

We make 100 simulations for these sample sizes and compare them with completely spatially random (CSR) samples of the same sizes in the XXL volume. In Section 3.3, we computed the nearest neighbours of the \( k = 8, 10, 12 \) and 14 order. Following the same procedure here, we obtain the nearest neighbour distances of the \( k = 12 \)th order for the XXL and the CSR samples and using the \texttt{ks.test()} procedure in \texttt{R}'s \texttt{stats} library, and compute the maximal difference between the cumulative distributions. We repeat this procedure between the CSR nearest neighbour distances. The distributions of KS differences between XXL–CSR and CSR–CSR samples are displayed in Fig. 7.

As one may infer from this figure, the distribution of KS differences between XXL–CSR samples does not differ significantly from those of CSR–CSR in the case of the sample sizes of \( N = 10000 \) and 5000. On the contrary, in the case of \( N = 10000 \) and 20 000 the difference between the XXL–CSR and CSR–CSR cases is very significant.

Based on this result, one may conclude that the simulated XXL samples with sizes of \( N = 10000 \) and 5000 do not differ from the CSR case. On the other hand, the samples with sizes of \( N = 10000 \) and 20 000 differ significantly from the CSR case. Each sample corresponds to some mean distance to the nearest object of the \( k = 12 \)th order. The computed mean distances are tabulated in Table 4.

| Sample size | Mean dist. (Mpc) | Prob. of CSR |
|-------------|-----------------|-------------|
| 1000        | 627             | 0.39        |
| 5000        | 351             | 0.71        |
| 10000       | 277             | <2.2e−16    |
| 20000       | 217             | <2.2e−16    |

At this point, however, it may be appropriate to repeat the remark made at the beginning of this subsection: the existence of the GRB ring cannot be explained by a simple density enhancement of the underlying baryonic and/or dark matter. It probably needs some coordinated star-forming activity among the responsible GRB hosts. In this case the spatial distribution of GRBs does not necessarily trace the underlying matter distribution in general. However, we cannot exclude the possibility that the XXL simulation does not correctly account for all possible large-scale structures and GRBs are mapping a structure that was not simulated.

Some concern may arise, however, concerning the interpretation of the ring as a true physical structure, and of the causal relationship between the GRBs displaying it. Suhtonenko et al. (2011) has pointed out that cosmic structures greater than 140 Mpc in comoving coordinates did not communicate with one another during the late stage of universal expansion preceding recombination. The skeleton of the web was created during the inflationary period (Kofman & Shandarin 1988) and evolved slowly following this epoch.

The volume of the shell between 0.78 < \( z < 0.86 \) is 20.2 Gpc\(^3\) in the comoving frame. The corresponding volume is 2.1 Gpc\(^3\) for \( z = 0.2 \) in the case of the SDSS main sample and 14.4 Gpc\(^3\) for \( z = 0.4 \) for LRGs, respectively. The volume of the shell is about 10 000 times larger than the volume of a typical supercluster found by Liivamägi et al. (2012) in the SDSS data.

Since the cosmic web evolves slowly, the structure of the GRB ring should exhibit the same general characteristics as those displayed by superclusters defining the web. Comparing the estimated number of superclusters with the number of detected GRBs, it appears that every thousand superclusters has produced on average one measured GRB. GRBs are therefore very rare events superimposed on the web, and the small probability of GRB detection casts serious doubt on the nature of the GRB ring as a real physical structure.

Taking these distributional characteristics into account suggests that the Ring is probably not a real physical structure. Further studies will be needed to reveal whether or not the Ring structure could result from a low-frequency spatial harmonic of the large-scale matter density distribution and/or of universal star-forming activity.

5 SUMMARY AND CONCLUSION

Motivated by the recent discovery of Horváth et al. (2013) revealing a large Universal structure displayed by GRBs, we study the spatial distribution of these objects in the comoving frame. The advantage of this approach is that GRBs (which are short transients) have footprints in this frame that do not change with time.
We assume in this approach that, for the spatially homogenous and isotropic case, the joint observed distribution of the GRBs can be factored into two parts: one part depends on the angular coordinates while the other part is radial and depends on the redshift.

This assumption can be tested in two different ways. The first method is essentially that used by Horvath et. al. which compares the conditional probability of the GRB angular distribution at different $z$ values. The second method tests whether resampling the GRBs randomly makes any statistical changes in the distribution in the 3D comoving frame.

We estimate the spatial density of GRBs by searching the angular separations of the $k$th order nearest neighbours. For these computations, we use the $\text{knn.dist}(x,k)$ procedure in the FNN library of the R statistical package. To compromise between the large variance of estimated densities at small $k$ values and the smearing out of real small-scale structures at large $k$ values, we use the spatial densities obtained by taking $k = 8, 10, 12$ and $14$.

Resampling the redshift distribution 10 000 times and calculating the spatial densities from the samples obtained in this way, we obtain mean densities and their variances assuming the null hypothesis, i.e. that the factorization of the spatial distribution of GRBs is valid. Subtracting the mean value from the observed one and dividing by the standard deviation, we calculate the standardized values of the densities obtained from the nearest neighbour procedure.

KS tests revealed that the logarithmic densities obtained in this way follow a Gaussian distribution allowing us to get a logarithmic likelihood function as a sum of the squared logarithmic densities. Since the sum of squared Gaussian variables follows a $\chi^2$ distribution with $k$ degrees of freedom, by selecting objects in a certain statistical package. To compromise between the large variance of estimated densities at small $k$ values and the smearing out of real small-scale structures at large $k$ values, we use the spatial densities obtained by taking $k = 8, 10, 12$ and $14$.

Since the calculated logarithmic densities in the $k = 8, 10, 12, 14$ cases are strongly correlated pairwise, performing a PCA allows us to join the logarithmic densities in the first (the only significant) PC variable representing 91 per cent of the total variance. Computing $\chi^2$ values from this PC for $k = 8, 10, 12, 14$ degrees of freedom and plotting them as a function of the distance, we find a very pronounced peak at about 2800 Mpc corresponding to a significance of 99.95, 99.93, 99.96 and 99.97 per cent, depending on the degrees of freedom.

We plot the angular positions of the GRBs within the FWHM range around the distance of the $\chi^2$ peak. Examining these plots, we conclude the following:

(1) There is a ring consisting of nine GRBs having a mean angular diameter of 36° corresponding to 1720 Mpc in the comoving frame.
(2) The ring is located in the $0.78 < z < 0.86$ redshift range having a standard deviation of $\sigma_z = 0.025$, corresponding to a comoving distance range of 2672 $< d_L < 2866$ Mpc having a standard deviation of $\sigma_d = 65$ Mpc.
(3) If one interprets the ring as a real spatial structure, then the observer has to see it nearly face-on because of the small standard deviation of GRB distances around the object’s centre.
(4) The ring can be a projection of a spheroidal structure. Adopting this approach, one has to assume that each host galaxy has a period of $2.5 \times 10^8$ yr during which the GRB rate is enhanced.
(5) The mass of the object responsible for the observed ring is estimated to be in the range of $10^{17} - 10^{18} M_\odot$ if the true structure is a torus or $10^{15} - 10^{16} M_\odot$ in case of a spheroid.
(6) GRBs are very rare events superimposed on the cosmic web identified by superclusters. Because of this, the ring is probably not a real physical structure. Further studies are needed to reveal whether or not the Ring could have been produced by a low-frequency spatial harmonic of the large-scale matter density distribution and/or of universal star-forming activity.

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REFERENCES
Angulo R. E., Springel V., White S. D. M., Jenkins A., Baugh C. M., Frenk C. S., 2012, MNRAS, 426, 2046
Aragon-Calvo M. A., Szalay A. S., 2013, MNRAS, 428, 3409
Bagla J. S., Yadav J., Seshadri T. R., 2008, MNRAS, 390, 829
Bahcall N. A., Kulier A., 2014, MNRAS, 439, 2505
Balzas L. G., Meszaros A., Horvath I., 1998, A&A, 339, 1
Balázs L. G., Mészáros A., Horváth I., Vavrek R., 1999, A&AS, 138, 417
Briggs M. S., 1993, ApJ, 407, 126
Castro Cerón J. M., Michalowski M. J., Hjorth J., Malesani D., Gorosabel J., Watson D., Fynbo J. P. U., Morales Calderón M., 2010, ApJ, 721, 1919
Centrella J., Melott A. L., 1983, Nature, 305, 196
Clowes R. G., Iovino A., Shaver P., 1987, MNRAS, 227, 921
Clowes R. G., Campuzano L. E., Graham M. J., Söchtging I. K., 2012, MNRAS, 419, 556
Clowes R. G., Harris K. A., Raghunathan S., Campuzano L. E., Söchtging I. K., Graham M. J., 2013, MNRAS, 429, 2910
Davidzon I. et al., 2013, A&A, 558, A23
Doroshkevich A. G., Klypin A. A., 1988, MNRAS, 235, 865
Doroshkevich A., Tucker D. L., Allam S., Way M. J., 2004, A&AA, 418, 7
Driver S. P., Popescu C. C., Tuffs R. J., Liske J., Graham A. W., Allen P. D., de Propris R., 2007, MNRAS, 379, 1022
Einasto J., Gramm M., 1993, ApJ, 407, 443
Einasto J., Hoeveer M., Saar E., 1980, Nature, 283, 47
Einasto J. et al., 1997, Nature, 385, 139
Einasto J. et al., 2006, A&A, 459, L1
Einasto J. et al., 2011, A&A, 534, A128
Einasto M. et al., 2014, A&A, 568, A46
Ellis G. F. R., 1975, QJRAS, 16, 245
Frisch P., Einasto J., Einasto M., Freundling W., Fricke K. J., Gramm M., Saar V., Toomet O., 1995, A&AA, 296, 611
Gott J. R., III, Juric M., Schlegel D., Hoyle F., Vogele M., Tegmark M., Bahcall N., Brinkmann J., 2005, ApJ, 624, 463
Graham M. J., Clowes R. G., Campuzano L. E., 1995, MNRAS, 275, 790
Haberzetti L. et al., 2009, ApJ, 702, 506
Haines C. P., Campuzano L. E., Clowes R. G., 2004, A&A, 421, 157
Heinimäki P., Sahholonen I., Saar E., Einasto M., Einasto J., Virtanen H., 2005, preprint (arXive-prints)
Horváth I., Hakki M., Bagoly Z., 2013, in 7th Huntsville Gamma-Ray Burst Symposium, GRB 2013: Paper 33 in eConf Proceedings C1304143
Horváth I., Hakki M., Bagoly Z., 2014, A&A, 561, L12
Icke V., 1984, MNRAS, 206, 1P
Icke V., van de Weygaert R., 1991, QJRAS, 32, 85
Kim J., Park C., Rossi G., Lee S. M., Gott J. R., III, 2011, J. Korean Astron. Soc., 44, 217
Kitaura F.-S., 2013, MNRAS, 429, L84
Kitaura F.-S., 2014, preprint (arXive-prints)
Kofman L. A., Shandarin S. F., 1988, Nature, 334, 129
Komberg B. V., Lukash V. N., 1998, Gravitation Cosmol., 4, 119
Kormberg B. V., Kvatsovs A. V., Lukash V. N., 1994, A&AA, 286, L19
Kormberg B. V., Kvatsovs A. V., Lukash V. N., 1996, MNRAS, 282, 713
Le Fèvre O. et al., 2013, A&A, 559, A14
Le Fèvre O. et al., 2014, Messenger, 155, 33
In this appendix, we compare the projection of two simple spherical models: (a) a homogeneous sphere and (b) a spherical shell. It is not difficult to derive the probability density functions for these projections into a plane. If we denote the projected distance from the mean position of the group to one of the members by \( g \), then we can normalize each position relative to the maximum projection \( g_{\text{max}} \), so that the projections vary in the range of \([0,1]\). In case of a homogeneous sphere the projected probability density is given by

\[
f(g) = 3g\sqrt{1-g^2}.
\]  

(A1)

and in the case of a shell we get

\[
g(g) = \frac{g}{\sqrt{1-g^2}}. \tag{A2}
\]

The shapes of these functions are displayed in Fig. A1.

The figure demonstrates that the projections for a shell result in a significant enhancement of the points close to the maximal distance from the centre; this is not the case for a projected homogeneous sphere. We can compute the probability of measuring all nine points outside of the median distance. By definition, the median splits the distribution into two parts of equal probability. The value of the median distance for a projected homogeneous sphere is \( g_{\text{median}} = 0.61 \).

We can calculate the probability of measuring all nine points in the \( 0.61 < g < 1 \) regime. To calculate the probability of this case, we invoke the \texttt{binom.test()} procedure available in the \texttt{R} statistical package. The probability of finding an object outside the median distance is \( p = 0.5 \) by definition. The probability for finding all nine points outside the median is given by the binomial distribution.

Similarly, we can calculate the probability of having all nine objects in a region outside the \( g_{\text{median}} = 0.61 \) median, assuming that the true spatial distribution of the points is a shell. Integrating the \( g(g) \) probability density in the \([0.61,1]\) interval we get

\[
\int_{0.61}^{1} g(g) \, dg = \int_{0.61}^{1} \frac{g}{\sqrt{1-g^2}} \, dg = 0.7924. \tag{A3}
\]

Inserting this probability into the binomial test expression we get \( p = 0.2192 \).

Summarizing the results of the tests performed above, we infer that the probability of observing a ring-like structure from a projected 3D homogeneous density enhancement only by chance is \( p = 3.9 \times 10^{-3} \) while assuming a projected 3D shell it is much higher, \( p = 0.22 \). This gives us a good reason to believe that the ring is a result of a projected 3D shell-like regular pattern.

\[\text{Figure A1. Comparison of the projected probability densities of a homogeneous sphere and a shell. The median of the sphere is indicated.}\]