The Effects of Mergers on the Formation of Disc-Bulge Systems in Hierarchical Clustering Scenarios

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\textbf{ABSTRACT}

We study the effects of mergers on the structural properties of disc-like systems by using Smooth Particle Hydrodynamical (SPH) numerical simulations in hierarchical clustering scenarios. In order to assess the effects of mergers on the mass distributions we performed a bulge-disc decomposition of the projected surface density of the systems at different stages of the merger process. We assumed an exponential law for the disc component and the Sérsic law for the bulges. We found that simulated objects at $z = 0$ have bulge profiles with shape parameters $n \approx 1$, consistent with observational results of spiral galaxies. The complete sample of simulated objects at $z = 0$ and $z > 0$ shows that $n$ takes values in the range $n \approx 0.4 - 4$. We found that secular evolution tends to produce exponential bulge profiles, while the fusion of baryonic cores tends to increase the $n$ value and helps to generate the correlation between $B/D$ and $n$. We found no dependence on the relative mass of the colliding objects. Our results suggest that mergers, through secular evolution and fusions, could produce the transformation of galactic objects along the Hubble sequence by driving a morphological loop that might also depend on the properties of the central galactic potential wells, which are also affected by mergers.

\textbf{Key words:} methods: numerical - galaxies: evolution - galaxies: interactions - galaxies: fundamental parameters

\section{INTRODUCTION}

The origin of the Hubble sequence is still a controversial issue. In particular, spiral galaxies seem to evolve along this phenomenological classification although the physical mechanisms behind these morphological changes are not fully understood. The structural parameters, the gas abundances and the star formation activity vary along galaxies in the Hubble sequence in the sense that disc-dominated systems are also the more gaseous ones, experiencing on-going star formation. The large database of galaxies gathered in the last years have allowed to get more detailed information on the properties of different morphological type galaxies in the local Universe.

Observational results show that when a double exponential decomposition is applied to the surface luminosity density of late-type spirals, the scalelengths of the bulge and disc components seem to be restricted to a certain value $< r_b/r_d > \approx 0.10$ (Courteau, de Jong & Broeils 1996, hereafter CdJB96). This restriction in the range of possible scalelengths for spirals has been interpreted as a proof for secular evolution to be the responsible mechanism for bulge formation from an already in place disc structure. Recently, MacArthur, Courteau & Holtzman (2002, hereafter MCH02) studied a larger sample of late-type spirals and carried out a bulge-disc decomposition assuming a Sérsic law for the bulge surface luminosity density. These authors found a continuous distribution of the shape parameter of the bulge ($n \approx 0.2 - 2$) with a maximum at $n \approx 0.90$. They also found a restricted range of values for the scalelengths of the bulge and disc when the shape parameter is $n \approx 1$, $< r_b/r_d > = 0.13 \pm 0.06$ in the $R$-band. Previous works have also found that bulges of spirals can be fitted by a Sérsic law with different shape parameters (e.g., Andredakis, Peletier & Ballcels 1995, hereafter APB95; Khosroshahi, Wadadekar & Kembhavi 2000, hereafter KWK00; Graham & de Block 2001, hereafter GdB01). APB95 and GdB01 also found a correlation between the luminosity bulge-to-disc ratio $B/D$ and the shape parameter of the bulge, which also correlates with morphological type. These observational results suggest a possible connection between the formation of the bulge and...
Several theories have been developed to explain the formation of the disc and bulge components. In the case of the disc systems, the standard model is based on three hypotheses: the angular momentum is acquired through cosmological torques (Peebles 1969), baryons and dark matter have the same specific angular momentum content \(J\), and this specific angular momentum is conserved during the collapse and cooling of baryons (Fall & Efstathiou 1980, hereafter FE80). This model has been successful in reproducing several observational results in analytical and semi-analytical models (e.g., Dalcanton, Spergel & Summers 1997; Mo, Mao & White 1998). However, serious problems arose in numerical simulations where an important angular momentum transfer from baryons to the dark matter haloes during mergers was detected, breaking the condition of \(J\) conservation. Domínguez-Tenreiro, Tissera & Sáiz (1998, hereafter DT98) showed that a disc-like structure with observational counterpart (Sáiz et al. 2001, hereafter S01) can be built up if a compact stellar bulge is allowed to form without depleting the gas reservoir of the system. These stellar bulges provide stability to the gaseous disc systems which are capable of conserving a non-negligible fraction of its angular momentum during violent events. Supernova energy feedback could also contribute to the formation of the disc component, regulating the star formation rate and preventing early catastrophic depletion of the gas into stars. Probably both mechanisms, the formation of a compact stellar bulge which assures the axisymmetrical character of the potential well and energy feedback, work together in nature to allow the formation of spiral galaxies (e.g. Weil, Eke & Efstathiou 1998).

The formation of bulges is a more complex task since several mechanisms could be acting together such as monolithic collapse (Gilmore & Wyse 1998), mergers (Kauffman, Guiderdoni & White 1994) and secular evolution (Pfenniger & Norman 1990). In general, analytical models and pre-prepared simulations have focused on one or two of them at the time (e.g., van den Bosch 1999; Aguerri, Balcells & Norman 1990). The drawback of our approach is a lower numerical resolution compared to those used in studies of pre-prepared mergers. Assessment of possible numerical problems are discussed throughout the paper.

In section 2 we present the analysis of the simulations. In section 3 we discuss the results. Section 4 summarizes the conclusions.

## 2 MODELS AND ANALYSIS

We run cosmological SPH simulations which include hydrodynamics and star formation (Tissera, Lambas & Abadi 1997). The simulated boxes represent typical regions of a standard Cold Dark Matter (CDM) universe of \(5 h^{-1}\) Mpc side and \(64^3\) particles \((\Omega = 1, \Lambda = 0, h = 0.5)\). We assume a baryonic density parameter of \(\Omega_b = 0.10\). All baryonic and dark matter particles have the same mass, \(M_p = 2.6 \times 10^8 M_\odot\). We have used a gravitational softening of 3.0 kpc and a minimum hydrodynamical smoothing length of 1.5 kpc. We performed three simulations with different realizations of the power spectrum (cluster normalized) and the same cosmological and star formation parameters. The simulated volume is the result of a compromise between the need to have a well-represented galaxy sample and enough numerical resolution to study the astrophysical properties of...
the simulated galaxies. We are confident that since we focus our analysis on small scale processes such as mergers and interactions, scale fluctuations of the order of the box size will have no significant effect on such local processes. Note also that we are interested in the effects that mergers have on the mass distributions of the galactic objects analysing them as individual events and not in connection with their evolution or environment.

The star formation algorithm used in these models is based on the Schmidt law (see Tissera et al. (1997) for details). Cold gas particles are eligible to form stars if they are denser than a certain critical value and satisfy the Jeans instability criterium (Navarro & White 1994). Gas particles are checked to satisfy these conditions at all time-steps of integration. Hence, as the gas cools down and is gathered at the centre of dark matter haloes, it is gradually transformed into stars according to the particular history of evolution of each galactic object. Only one free parameter has to be fixed: the star formation efficiency, which links the dynamical time of the gas cloud represented by a particle with its star formation timescale. We have used the value adopted by S01 since it is adequate to reproduce disc-like structures with observational counterparts at \( z = 0 \). These authors used a low star formation efficiency which allows the formation of a stellar bulge that assures the axisymmetrical character of the potential well but without exhausting the gas reservoir (see also DT98). As a consequence, disc-like structures can be formed, although they remain mainly gaseous. Conversely, observed discs are mostly stellar, but they inherit the structural and dynamical properties of the gaseous discs out of which disc stars are formed. Hence, for the sake of comparison, the values of the parameters describing those properties can be reliably estimated from the simulated gaseous discs. In this paper, we will determine and use these parameters to study the effects of mergers on the mass distribution of the simulated galactic objects.

At \( z = 0 \), we identify galactic objects at their virial radius, analysing only those with more than 4000 dark matter particles within their virial radius in order to diminish numerical resolution problems. Within each galactic object, a main baryonic clump is individualized which will be, hereafter, called the galaxy-like object (GLO). The selected GLOs have very well-resolved dark matter haloes which provide adequate potential wells for baryons to collapse in. This fact assures a reliable description of the gas density profiles (see Steinmetz & White 1997), which allows us to follow the star formation history of the GLOs (see Tissera 2000). With this strong restriction on the minimum number of particles, the final GLO sample at \( z = 0 \) is made up of 12 GLOs with virial velocities in the range 140-180 km/s.

We follow the evolution of the selected GLOs with lookback time (\( \tau (z) = 1 - (1 + z)^{-3/2} \)) identifying mergers and starbursts (SBs). We then construct their mergers trees and star formation rate (SFR) histories (because of the restrictive SF efficiency used, these SFR histories are mainly those of the bulge components). During a merger event, the progenitor object is chosen as the more massive baryonic clump within this merger tree, while the minor colliding baryonic clump is referred to as the satellite. Following T02 we study only those mergers that are directly linked to starbursts in the SF history of each GLO. These authors found that during the orbital decay phase of some merger events, early gas inflows which trigger star formation can develop. In this case, if the gas reservoir of the GLOs is not exhausted during these starbursts, second bursts are induced at the actual physical contact of the baryonic clumps. The induction of early gas inflows could be directly linked to the properties of the total potential wells of the systems in agreement with previous results reported by Barnes & Hernquist (1996) and Mihos & Hernquist (1996) in pre-prepared simulations. When no gas inflows are triggered during the orbital decay phase (ODP), only one SB is detected when the two baryonic clumps collide. These events will be called single SBs (SSBs), while those where two SBs are detected will be called double SBs (DSBs). We studied a total of 18 merger events with merger parameters settled by the cosmological model. Among them, 11 are classified as DSBs and 7 as SSBs. Fig. shows an example of both types of SBs. We have also plotted the distance between the centres of mass of the progenitor and the satellite since the redshift they share the same dark matter halo. The merger event will be determined by four redshifts of reference, \( z_A \): the beginning of the first bursts in DSBs and the ODP in SSBs, \( z_B \): the end of both the first bursts in DSBs and the ODP in SSBs, \( z_C \): the beginning of the second SBs in DSBs and the only bursts in SSBs (in this case \( z_B = z_C \)), and \( z_D \): the end of the second bursts in DSBs and of the only one in SSBs. These redshifts have been depicted in Fig. Note that the ODP, defined as the time since two objects share the same dark matter halo to the fusion of their main baryonic cores, is determined by \( z_A \) and \( z_C \). The \( z_B \) redshift is used in the case of DSBs since it establishes the end of the first SBs.

In order to quantify the changes in the baryonic mass distributions, we carried out a bulge-disc decomposition of the projected mass surface density of the GLOs at \( z = 0 \) and of the progenitors at each of the redshifts that define a merger event. We assumed an exponential profile for the potential well but without exhausting the gas reservoir (see also DT98). As a consequence, disc-like structures can be formed, although they remain mainly gaseous. Conversely, observed discs are mostly stellar, but they inherit the structural and dynamical properties of the gaseous discs out of which disc stars are formed. Hence, for the sake of comparison, the values of the parameters describing those properties can be reliably estimated from the simulated gaseous discs. In this paper, we will determine and use these parameters to study the effects of mergers on the mass distribution of the simulated galactic objects.

Bulge-disc connection

In order to quantify the changes in the baryonic mass distributions, we carried out a bulge-disc decomposition of the projected mass surface density of the GLOs at \( z = 0 \) and of the progenitors at each of the redshifts that define a merger event. We assumed an exponential profile for the disc component and the Sérsic law for the bulge central mass concentration:

\[
\Sigma_d = \Sigma_0^d \exp\left(-\frac{(r/r_d)^{1/n}}{1}\right) \\
\Sigma_b = \Sigma_0^b \exp\left(-\frac{(r/r_b)^{1/n}}{1}\right)
\]

where \( r_d \) and \( r_b \) are the disc and bulge scalelengths, \( \Sigma_0^d \) and \( \Sigma_0^b \) the corresponding central mass surface densities and \( n \) the bulge shape parameter.

For each GLO at \( z = 0 \) and its progenitor objects, we estimated the total angular momentum \( (j) \) and projected the mass distribution onto the perpendicular plane defined by \( j \). We then integrated the projected baryonic mass in concentric cylinders of radius \( r \) since, as it has been proved by S01, it is numerically more convenient to fit the integrated mass than the surface density. In order to diminish the effects of numerical resolution, the fits were made from \( r = 1.5 \) kpc to \( r \approx 30 \) kpc, except in those cases where the mass distributions of the progenitors were strongly perturbed by the entrance of the satellites.

The total integrated baryonic mass of the GLOs (\( M_{GLO} \)) is estimated, on one hand, as the baryonic mass within \( r = 30 \) kpc, and will be associated with the luminous galaxy. On the other hand, \( M_{GLO} \) is the result of the disc (\( M_d \)) and bulge (\( M_b \)) mass contributions:

\[
M_{GLO} = M_d + M_b
\]
late-type spiral galaxies in the K-band (27 early-type disc galaxies in the K-band), APB95 (30 S0-Sbc galaxies in deep U-band), CdJB96 (173 spiral Sb-Sc galaxies in deep U-band), KWK00 (n ≤ 1.5: open squares), KWK00 (n ≤ 1.5: open triangles), MCH02 (n ≤ 1.5: open circles) are also included.

3 RESULTS

3.1 GLOs at z = 0

As we mentioned before, S01 showed that disc-like objects formed in numerical simulations where stellar bulges were allowed to form have structural parameters and specific angular momentum content similar to those of observed spiral galaxies. For the sake of completeness and to show that these findings are valid for our GLOs, we will first discuss their properties at z = 0. The structural parameters and the circular velocity ($V^2 = GM(r)/r$) at $r = 2.2r_d$ allow the classification of our objects as intermediate spirals ($1 < n < 2$), normal spirals ($2 < n < 3$), spiral ($n > 3$), and compact bright spiral ($n > 5$). This is the result of our strong condition on the mass of GLOs in order to work with those better resolved.

In Fig. 3 we show the relation between the disc and bulge scalelengths. Each GLO has a label that specifies its n shape parameter. We include the observational data from CdJB96, MGH99, KWK00 and MCH02. Note that for increasing n parameters, the scalelengths tend to move to the upper envelope of the distribution. This effect is also seen in the observed scalelength distribution. We have adopted $n = 1.5$ as a general cut-off value to segregate the samples between those with exponential and non-exponential profiles for the bulge components.

In Fig. 4 we show the normalized histogram of n for GLOs at z = 0 (a), and the corresponding distribution for the observations of APB95, KWK00 and MCH02 (b), normalized to the total number of galaxies in each sample.
Figure 3. Histogram of the shape parameter of the bulges for simulated objects at $z = 0$ (a) and observations from APB95 (dashed line), KWK00 (dotted line) and MCH02 (solid line), normalized to the total number of galaxies in each sample.

The combined observational sample covers a morphological range from Sa to Sc galaxies, but this is not a consistent sample. We show it in order to broadly compare the range of $n$ values obtained from the simulations with observations. The average $n$ values over the observed subsamples are $<n_{\text{APB95}}> = 2.97 \pm 1.29$, $<n_{\text{KWK00}}> = 2.88 \pm 0.94$ and $<n_{\text{MCH02}}> = 1.02 \pm 0.43$ while the simulated sample has $<n_{\text{sim}}> = 1.23 \pm 0.40$. The simulated distribution is consistent with the observations of MCH02 at one $\sigma$-level, which suggests that our objects are more similar to late-type galaxies. However, $<n_{\text{sim}}> >$ value also agrees at 3 $\sigma$-level with the corresponding mean values of the early-type samples.

A key point in the formation of discs is the conservation of the specific angular momentum (FE80). Our SF algorithm has been implemented in such a way that it has allowed the formation of stellar bulges but without exhausting the gas reservoir (DT98). Hence, gaseous discs have been able to regenerate after mergers. In order to compare the angular momentum ($j$) of the three mass components: bulge, disc and dark matter halo, with observations, we plot in Fig. 4 their specific angular momentum ($J = j/M$) content as a function of mass ($M$). The mass and angular momentum of bulges have been estimated with stars within $r < r_{\text{eff}}$, while those of the gaseous discs have been calculated within $r < 3.2 r_{\text{d}}$. These criteria have been adopted following those used by observers. The mass and angular momentum of the dark matter haloes have been estimated at the virial radius. In this figure, we have also included two boxes that depict the observational region covered by spirals and ellipticals as given by Fall (1983). From this plot we can see that the disc components and the dark matter haloes have comparable specific angular momentum contents as predicted by the standard disc formation model of FE80. Conversely, stellar bulges have been formed from material that has lost most of its angular momentum.

From these results we conclude that the simulated GLOs have structural and dynamical parameters that statistically resemble those of current normal spirals at $z = 0$. In the following section we will study the role played by mergers in the determination of these properties and try to understand the origin of correlations among the structural parameters such as $B/D$ versus $n$.

3.2 Analysis of progenitors during merger events

We carried out the bulge-disc (B+D) decomposition of the progenitor objects of the GLOs analysed in the previous section during merger events at the four redshifts of interest: $z_A$, $z_B$, $z_C$ and $z_D$. In this section we will study how their structural parameters change during these processes.
Table 1. Mean values of bulge scalelength ($r_b$), disc scalelength ($r_d$) and bulge shape parameter ($<n>$) for single (S) and double (D) SBs for the redshifts of interest.

| $z_i$ | $r_b$ [kpc] | $r_d$ [kpc] | $<n>$ |
|-------|-------------|-------------|-------|
| $z_A$ S | 0.46±0.17   | 6.2±0.7     | 1.5±0.2 |
|        D | 0.31±0.11   | 4.9±0.6     | 1.8±0.2 |
| $z_B$ S | 0.57±0.16   | 5.6±0.7     | 1.3±0.2 |
|        D | 0.37±0.16   | 5.9±0.4     | 1.6±0.2 |
| $z_C$ S | 0.57±0.16   | 5.6±0.7     | 1.3±0.2 |
|        D | 0.53±0.12   | 5.0±0.4     | 1.2±0.1 |
| $z_D$ S | 0.52±0.12   | 5.0±0.7     | 1.4±0.3 |
|        D | 0.37±0.08   | 5.4±0.6     | 1.4±0.2 |

Note: The errors correspond to the bootstrap method, $\sigma_{bi}$.

3.2.1 The shape parameter

In Fig. 5(a) we show the $r_d$ versus $r_b$ relation for all progenitors during merger events and at $z = 0$. We have also included the observational data from a B+D decomposition from MGH99, KWK00 and MCH02, and from a double exponential decomposition by CDJB96. The combined observed sample of spirals shows a certain correlation between $r_d$ and $r_b$, although some of them have, for a given disc scalelength, a smaller bulge one than that expected from this relation. We also note that the scalelengths of our simulated disc structures show a similar behaviour. In order to individualize which observed and simulated objects belong to each of these two distributions of scalelengths, we segregate them according to their shape parameters as shown from Fig. 5(b) to Fig. 5(d). It is clear that as galaxies and simulated objects with large shape parameters are taken out, the relation for both, observed and simulated scalelengths, gets better defined.

We also note that the $r_b$ versus $r_d$ relation for GLOs and galaxies with bulges with $n > 1.5$ shows no trend at all. It is encouraging that the observed and simulated disc structures show scalelength distributions with similar patterns, for both $n \approx 1$ or $n \neq 1$ bulge profiles. In the case of exponential bulges (0.9 $\leq n \leq 1.1$), the correlation is well defined with $<r_b/r_d> = 0.14 \pm 0.05$ and $<r_b/r_d> = 0.13 \pm 0.06$ for observations and simulations, respectively. Table 2 summarizes the mean values of $r_b$, $r_d$ and $n$ for GLOs at the four $z$ of interest.

Similary to Fig. 5(a) we display in Fig. 6 the histogram of shape parameters for all progenitors at all studied redshifts. Hence this distribution gathers the information on the $n$ parameters of different progenitors at different stages of evolution. From the comparison with that shown in Fig. 5 corresponding to the bulges of the final structures at $z = 0$, we deduce that in the past of current GLOs, their progenitors went through stages where their bulges had $n > 1.5$ parameters ($\sim 35$ per cent of the total sample). Meanwhile, GLOs at $z = 0$ show a trend to have $n \approx 1$, with less than 10 per cent being described by $n > 1.5$.

In order to assess the effects of mergers on the shape parameters of the objects and to look for possible differences between the effects of secular evolution and fusions, we calculate the percentages corresponding to the increase and decrease of the shape parameter during the intervals $z_A$-$z_C$, $z_C$-$z_D$ and $z_A$-$z_D$ for SSBs and DSBs. The results are shown in Table 2. Note that we are taking the interval $z_A$-$z_C$ as the total period corresponding to the ODP. For both single and double SBs, we find that bulges tend to decrease their $n$ parameters during the ODP, and that this effect is more important for DSBs. In contrast, the actual fusion of the clumps ($z_C$-$z_D$) tends to make the $n$ value of the bulges larger.

Therefore we find that the physical encounter of the baryonic cores tends to produce bulges with larger shape parameters (i.e. more concentrated), and $n \approx 1$ profiles tend to be formed by secular evolution. These results together with the fact that at $z = 0$ most of GLOs show $n \leq 1.5$, regardless of their past history, suggest that successive mergers have driven the bulges from exponential to more concentrated profiles, to come back to exponential ones, if there is available gas in the systems.

We think that the possibility of having secular evolution and core fusion associated to a merger helps to drive this cycle in the shape parameter. Note that the collisionless merger of stellar discs will not produce the same distributions. As a matter of fact, in order to detect this morphological loop, simulations should include gas dynamics and star formation.

![Figure 6. Total distribution of bulge shape parameters for the progenitors during merger events.](image-url)
Figure 5. Disc scalelength ($r_d$) as a function of bulge scalelength ($r_b$) for simulated objects at $z = 0$ and progenitors (filled pentagons) for which bulges are described by $0 \leq n \leq 5$ (a), $0.5 \leq n \leq 2$ (b), $0.8 \leq n \leq 1.5$ (c) and $0.9 \leq n \leq 1.1$ (d). Observational data from CdJ96 (crosses), MGH99 (open squares), KWK00 (open triangles) and MCH02 (open circles) for the same shape parameter intervals are also shown.

since, otherwise, bulges can not be re-shaped back to $n \approx 1$ (A01).

Fig. 7 shows the variations of the $n$ parameter during the ODP (a) and those produced during the fusion of the baryonic clumps (b) as a function of $M_{\text{star}}^z/M_{\text{star}}^0$, where $M_{\text{star}}^z$ is the stellar mass formed in the progenitors at $z$ and $M_{\text{star}}^0$ is the total stellar mass of the GLOs at $z = 0$. This ratio is an estimate of the presence of a stellar bulge since in these simulations the stars form preferentially in the dense regions. T02 found that the properties of the potential well can be directly linked to the stability of a system, in the sense that the shallower the potential well, the more susceptible the GLO to experience early gas inflows (i.e. secular evolution). The formation of stellar bulges has been found to inhibit such early gas inflows. This picture is in agreement with our results, where the major changes in the shape parameter during a merger event correspond to the GLOs that experience DSBs ($M_{\text{star}}^z/M_{\text{star}}^0 \leq 0.45$), which are less stable objects and can be strongly perturbed by the collision with a satellite. In contrast, more stable systems ($M_{\text{star}}^z/M_{\text{star}}^0 > 0.45$) would not experience such important perturbations in its mass concentration during the ODP (Table I).

We can also see from Fig. 7(a) that the most of the changes in $n$ are negative and can be linked to smaller $M_{\text{star}}^z/M_{\text{star}}^0$ values ($M_{\text{star}}^z/M_{\text{star}}^0 \leq 0.45$). This could be related to the fact that stellar bulges can provide stability to the systems, making them less vulnerable to the influence of a satellite, principally during the ODP. From Fig. 7(b) we see that the fusion of the baryonic cores produce more positive changes in the $n$ parameters and hence, a trend to increase it. Again, the larger variations are related to smaller bulges (i.e. smaller $M_{\text{star}}^z/M_{\text{star}}^0$).

We also analysed the possibility that the changes in the $n$ value during mergers are linked to the relative mass of the colliding systems, finding no correlation in contrast to the results of A01. However, these authors only considered collisionless mergers while in all our mergers dissipation plays an important role.

In Fig. 8 we have plotted the averaged $n$ parameters of the progenitors at different $z$ intervals. We can see that, although the bootstrap errors are high, there is a trend for the averaged $n$ to decrease toward $n = 1$ for decreasing $z$. 

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Figure 7. Variations in the shape parameter of the progenitors during the orbital decay phase (a) and during the fusion of the baryonic clumps (b) in single (filled circles) and double (filled squares) SBs, as a function of the fraction of stars already formed in the progenitors, \( M_{\text{star}}^z/M_{\text{star}}^{z=0} \) when the mergers start. \( M_{\text{star}}^z/M_{\text{star}}^{z=0} = 0.45 \) (dashed lines) has been taken to be an indicator of the presence of a well-formed stellar bulge after T02.

Figure 8. Average \( n \) parameters of the progenitors at different \( z \) intervals and of the GLOs at \( z = 0 \). The error bars correspond to the bootstrap method.

Nevertheless, at all redshifts we found the mean \( n \) smaller than \( n = 2 \), indicating that, in spite of the spread of the parameters, the majority of the bulges are best fitted with \( n \lesssim 1.5 \).

3.3 The \( r_b/r_d \) ratio

It has been argued by some authors (e.g., CdJB96) that there is a restricted range of \( r_b/r_d \) for spiral galaxies which has been interpreted as the result of secular evolution of the central mass concentrations. In Fig. 8 we showed the \( r_d \) versus \( r_b \) distribution for GLOs at \( z = 0 \) and for observations. We have combined the observational results from MGH99, KWK00 and MCH02 which include spiral galaxies from Sa to Sc. For these distributions we find \( < r_b/r_d > = 0.07 \pm 0.02 \) (the bootstrap error is \( \sigma_{\text{bt}} = 0.02 \)) and \( < r_b/r_d > = 0.10 \pm 0.08 \) (\( \sigma_{\text{bt}} = 0.01 \)) for the simulated and observed data, respectively, showing a good agreement.
However, we have already shown in Fig. 3 that the simulated and observed scalelength distributions show a clear trend only for those objects with bulges with $0.9 < n < 1.1$. Hence, we divided the observations and simulated data into three subsamples according to their $n$ parameters ($n < 0.9$, $0.9 < n < 1.1$ and $n > 1.1$) and estimate their mean $r_b/r_d$. Table 3 summarizes the results. First, we note that observational and simulated $<r_b/r_d>$ in the three subsamples are in very good agreement, supporting the hypothesis that GLOs in CDM models have observational counterpart. We also see that the larger dispersion is found for bulges with $n < 0.9$ in both the observed and simulated data, and that these subsamples show higher $<r_b/r_d>$ than the subsamples with $n \approx 1$, suggesting that these systems systematically have larger $r_b$ (for this subsample: $<r_b> = 0.89 \pm 0.10$ and $<r_d> = 4.77 \pm 0.40$).

On the other hand, we find that systems with bulges better fitted by $n > 1.1$ show the smallest $<r_b/r_d>$. Hence for a given disc scalelength, the larger the concentration of the bulges, the smaller its scalelength (for this subsample: $<r_b> = 0.23 \pm 0.03$ and $<r_d> = 5.52 \pm 0.30$). If we assume that the $n$ parameter correlates with morphology as it is indicated by several works (e.g., de Jong 1996; Graham & Prieto 1999, hereafter GP99; GdB01), then the scalelength ratios should correlate with $n$. Fig. 4 shows the $r_b/r_d$ ratio as a function of the shape parameter for the simulated objects during mergers and for $z = 0$. We have also included observational data from KWK00 and MCH02. We find an anticorrelation between these two parameters, in agreement with observations. Assuming that the shape parameter is an indicator of Hubble type, this anticorrelation indicates that late-type spirals ($n \leq 1.5$) have larger $r_b/r_d$ ratios than early-type spirals ($n > 1.5$). We also note that the relation for the observed and simulated ratios seems to flatten for $n \ll 1$. A simple extrapolation gives a value of $r_b/r_d \approx 0.40$.

Because the shape parameter and the scalelength of the Sérsic law are coupled parameters, this trend of $r_b/r_d$ with $n$ could be the result of the fitting formula. However, Trujillo, Graham & Caon (2001, hereafter TGC01) argued that there is a physical trend stronger than that resulting from the fitting formula. Because we use the same expression for all fittings, the fact that there is a segregation in the scalelength distribution according to the value of $n$ suggests that this trend could have physical basis and could be related to a difference in the evolution of the GLOs or in the efficiency of some physical process. This is confirmed when $r_{eff}$ is used instead of $r_b$ for which we find $<r_{eff}/r_d> = 0.29$ ($\sigma_B = 0.04$) for $n < 0.9$, $<r_{eff}/r_d> = 0.26$ ($\sigma_B = 0.03$) for $0.9 < n \leq 1.1$, and $<r_{eff}/r_d> = 0.30$ ($\sigma_B = 0.05$) for $n > 1.1$ (the combined observational sample of KWK00 and MCH02 shows $<r_{eff}/r_d> = 0.23$ ($\sigma_B = 0.02$) for $n < 0.9$, $<r_{eff}/r_d> = 0.28$ ($\sigma_B = 0.04$) for $0.9 < n \leq 1.1$, and $<r_{eff}/r_d> = 0.35$ ($\sigma_B = 0.03$) for $n > 1.1$). We have used the approximation of Ciotti & Bertin (1999) to transform simulated $r_b$ into $r_{eff}$.

Let us now look at the mean $<r_b/r_d>$ during mergers. In Table 3 we show the mean $r_b/r_d$ at $z_A$ and $z_B$ for SSBs and DSBs. We find no differences between the mean $r_b/r_d$ but both samples agree in increasing the mean ratios and the dispersions at $z_D$. Fig. 10 shows the distributions of scalelength during mergers. Note that, independently of the stage of evolution of the mergers, the distributions have similar characteristics and are within the observed range at $z = 0$. Finally in Fig. 3 in the small box, we show how the averaged $r_b/r_d$ moves throughout mergers associated with single (thin line) and double (thick line) starbursts. It is clear that early gas inflows during the ODP tend to decrease $n$ in both cases but for DSBs the changes are more important, as expected from their stability properties (see Fig. 2 and T02). Conversely, the fusion of the cores moves the distribution toward higher $n$.

### 3.4 The B/D ratio

The luminosity $B/D$ ratio has been traditionally used as an indicator of morphology (e.g., Andredakis & Sanders 1994). The correlation between $B/D$ and $n$ has been shown by APB95 and recently confirmed by TGC01. Linear regression through observations show $\partial(B/D)/\partial n = 1.0 \pm 0.3$ ($\sigma_B = 0.3$) for APB95, $0.8 \pm 0.5$ ($\sigma_B = 0.5$) for KWK00 and $1.7 \pm 0.3$ ($\sigma_B = 0.3$) for MCH02. The combined sample has $\partial(B/D)/\partial n = 1.2 \pm 0.1$ ($\sigma_B = 0.1$).

Note that in the models we have masses instead of luminosities and that the mass-to-light ratios of the discs ($\gamma_d$) and bulges ($\gamma_b$) could be different because of their different stellar populations. However, their correct values and/or possible dependence on radius as well as redshift are still unclear. In order to match the luminosity $B/D$ we have to rescale the mass $B/D$ ratio by a constant factor of 20. We address that we are using the same mass-to-light ratios for all GLOs regardless of the redshift[4] disc scalelength or observational band.

In Fig. 11 we show the $B/D$ ratio for the simulated GLOs during merger events at the four redshifts of interest and the observational data from APB95, KWK00 and MCH02. The low statistical number restrict the use of linear regression through the simulated data. However, we see

| $z$ | $<r_b/r_d>$ | $<r_b/r_d>$ | $<r_b/r_d>$ |
|-----|-------------|-------------|-------------|
| $z \neq 0$ | $0.10 \pm 0.01$ (0.01) | $0.07 \pm 0.02$ (0.02) | $0.10 \pm 0.08$ (0.008) |
| $z = 0$ | $0.24 \pm 0.20$ (0.05) | $0.16 \pm 0.07$ (0.01) | $0.13 \pm 0.06$ (0.02) |
| $n < 0.9$ | $0.05 \pm 0.05$ (0.01) | $0.03 \pm 0.04$ (0.01) | $0.07 \pm 0.05$ (0.02) |
| $n \leq 0$ | $0.07 \pm 0.08$ (0.03) | $0.11 \pm 0.08$ (0.03) | $0.10 \pm 0.10$ (0.04) |

1 Nevertheless, we found that very similar $\gamma_b/\gamma_d$ ratios are needed to match the observed luminosity $B/D$ for progenitors at different $z$, although our dispersion is quite high. Note that we have used the total baryonic masses to estimate luminosities. In a more realistic model including energy feedback, some fraction of the cold gas would be reheated. Hence, the values required for our models to match the observed luminosities could be taken as upper limits to the mass-to-light ratios.
Figure 10. Distribution of bulge and disc scalelenghts for the simulated galactic objects at $z_A$ (a), $z_B$ (b), $z_C$ (c) and $z_D$ (d). Different symbols have been used to distinguish between $n \leq 1.5$ (encircled symbols) and $n > 1.5$, and between double (filled squares) and single (filled circles) starbursts. Observational data from MGH99 ($n = 1$ crossed squares, $n = 2$: open squares), KWK00 ($n \leq 1.5$: crossed-triangles, $n > 1.5$: open triangles) and MCH02 ($n \leq 1.5$: crossed-circles, $n > 1.5$: open circles) are also shown.

that, at any time during merger events, the simulated values are within the observational range.

In Fig.12, we plot the variations $n_{z_i} - n_{z_j}$ versus $B/D_{z_i} - B/D_{z_j}$ where $i/j$ can be A, C or D. From this figure we see that GLOs which experience gas inflows show more important changes in both $n$ and $B/D$ during the ODP. In the case of those GLOs experiencing SSBs (e.g., no early gas inflows) there is no important change in the shape parameter or the $B/D$ ratios during the ODP. We also note that the changes in $n$ and $B/D$ do not correlate. Conversely, during the actual fusion of the objects (from $z_C$ to $z_D$) both types of GLOs experience large changes in both parameters in the expected direction: the larger the increase in the shape parameter, the larger the corresponding increase of the $B/D$ ratio or vice-versa. Fig.12(c) shows the overall changes over the merger events. This distribution has higher dispersion than that obtained during the fusions because it also combines the effects of secular evolution. However, the overall effects of mergers is to settle the correlation between $B/D$ and $n$ so that more concentrated bulges (higher shape parameters) have the larger $B/D$ ratios. However, we note that, although the range of values can be matched, the simulated $B/D$ do not show the same correlation signal ($r = 0.20$) than the observations ($r = 0.55$).

4 CONCLUSIONS

We have studied the properties of the baryon distributions in galaxy-like objects focusing on the effects of mergers. We resort to observations of spiral galaxies of different morphology to constrain our findings.

We found that on average, galactic objects formed in hierarchical clustering scenarios reproduce the angular momentum and structural parameter distributions of spiral galaxies, if a stellar bulge is allowed to form and early gas depletion is avoided. We have succeeded in these two aspects but at the expense of inhibiting star formation on discs. A consistent treatment of energy feedback may help to remove this caveat.

These simulations have allowed us to study how mergers change the distribution of baryons by analysing the evolution of their structural parameters. We found that, on average, galactic objects tend to have nearly exponential bulges...
at all redshift. However, note that these simulations produce bulges with shape parameters in the range 0.5 – 4.

For those systems with \( n \approx 1 \) bulges we found a correlation among their bulge and disc scalelengths in very good agreement with observations. However, for \( n > 1 \) the scalelength distribution is disorder and displaced to smaller bulge scalelengths. Observations show the same behaviour. The opposite distribution is found for systems with \( n < 1 \) bulges which have larger \( r_b/r_d \) values and larger dispersions. We found the disc scalelengths to be approximately independent of \( n \) parameters, so that this correlation implies that more concentrated objects have smaller bulge scalelengths (or larger effective radius). We also found a maximum \( r_b/r_d \) of \( \approx 0.40 \) for \( n \rightarrow 0 \). Higher numerical resolution simulations are needed to study the formation of such low mass surface profiles.

In order to test our results for low resolution problems in the determination of the structural parameters of the galaxy-like objects, we followed Steinmetz & Müller (1994) and used the bootstrap technique to estimate the effects of low particle number statistics within the objects. We calculated an accuracy better than 25 per cent for the scalelengths and shape parameters.

We found that gas inflows during the orbital decay phase tend to produce important changes in the mass distributions generating \( n \approx 1 \) profiles, while the fusions of the baryonic cores tend to increase the \( n \) parameter. As a consequence, a morphological loop can be driven by mergers which might be responsible of triggering secular evolution as well as of the baryonic core fusions. The triggering of secular evolution is found to be linked to the presence of a stellar bulge so that systems with well-formed stellar bulges do not experience early gas inflows. Hence, the pace of this morphological loop could be regulated by the properties of the galactic central potential wells which are also affected by the merger history of the objects (see also Tissera & Domínguez-Tenreiro 1998 and Tissera 2000).

We found that the simulated mass bulge-to-disc ratios are within observed range. It is also noted that during the ODP larger changes are observed in the \( B/D \) ratios of those objects that experience gas inflows. The changes during this period are, however, quite disorder. It is at the fusion of the baryonic clumps that changes in the shape parameters are correlated with changes in the bulge-to-disc ratio. In our simulations, the actual fusions are responsible of significantly increasing the mass concentration at the centre.
Figure 12. Variations in the shape parameters of the progenitors during the orbital decay phase (a), the fusion of the baryonic clumps (b) and the whole merger process (c) in single (filled circles) and double (filled squares) SBs as a function of the corresponding changes in the bulge-to-disc ratio, $B/D$.

Hence, we found that the fusion of the baryonic cores could be the process that determine the observed correlation between the luminosity $B/D$ ratio and the shape parameter or morphological type. However, we found no dependence on the relative masses of the colliding objects.

Overall, our results indicate that the morphological properties of galactic objects are the result of their merger histories within a hierarchical clustering scenario. Based on the good agreement found so far with observations we support the hypothesis of mergers as the main morphological driver along the Hubble sequence. Consequently, the particular and detailed history of substructure aggregation could be a key point in the determination of the astrophysical properties of galaxies.

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