MASS-TO-LIGHT RATIOS OF FIELD EARLY-TYPE GALAXIES AT $z \sim 1$ FROM ULTRADEEP SPECTROSCOPY: EVIDENCE FOR MASS-DEPENDENT EVOLUTION

A. VAN DER WEL AND M. FRANX
Leiden Observatory, P.O. Box 9513, NL-2300 AA Leiden, Netherlands; vdwel@strw.leidenuniv.nl

P. G. VAN DOKKUM
Department of Astronomy, Yale University, P.O. Box 208101, New Haven, CT 06520-8101

H.-W. RIX
Max-Planck-Institut für Astronomie, Königstuhl 17, D-69117 Heidelberg, Germany

G. D. ILLINGWORTH
University of California Observatories/Lick Observatory, University of California, 373 Interdisciplinary Sciences, Santa Cruz, CA 95064

AND

P. ROSATI
European Southern Observatory, Karl-Schwarzschild-Strasse 2, D-85748 Garching, Germany

Received 2004 November 15; accepted 2005 March 24

ABSTRACT

We present an analysis of the fundamental plane for a sample of 27 field early-type galaxies in the redshift range $0.6 < z < 1.5$ in the Chandra Deep Field–South and the field of the background cluster RDCS 1252.9–2927. Sixteen of the galaxies are at $z > 0.95$. The galaxies in this sample have high signal-to-noise ratio spectra obtained at the Very Large Telescope and high-resolution imaging from the HST Advanced Camera for Surveys. From comparison with lower redshift data, we find that the mean evolution of the mass-to-light ratio ($M/L$) of our sample is $\Delta \ln (M/L_B) = (-1.74 \pm 0.16)z$, with a large galaxy-to-galaxy scatter. The strong correlation between $M/L$ and rest-frame color indicates that the observed scatter is not due to measurement errors, but due to intrinsic differences between the stellar populations of the galaxies, such that our results can be used as a calibration for converting luminosities of high-redshift galaxies into masses. This pace of evolution is much faster than the evolution of cluster galaxies. However, we find that the measured $M/L$ evolution strongly depends on galaxy mass. For galaxies with masses $M > 2 \times 10^{11} M_\odot$, we find no significant difference between the evolution of field and cluster galaxies: $\Delta \ln (M/L_B) = (-1.20 \pm 0.18)z$ for field galaxies and $\Delta \ln (M/L_B) = (-1.12 \pm 0.06)z$ for cluster galaxies. The relation between the measured $M/L$ evolution and mass is partially due to selection effects, as the galaxies are selected by luminosity, not mass. We calculate the magnitude of this effect for the subsample of galaxies with masses higher than $M = 6 \times 10^{10} M_\odot$: the uncorrected value of the evolution is $\Delta \ln (M/L_B) = (-1.54 \pm 0.16)z$, whereas the corrected value is $(-1.43 \pm 0.16)z$. However, even when taking selection effects into account, we still find a relation between $M/L$ evolution and mass, which is most likely caused by a lower mean age and a larger intrinsic scatter for low-mass galaxies. Results from lensing early-type galaxies, which are mass selected, show a very similar trend with mass. This, combined with our findings, provides evidence for downsizing, i.e., for the proposition that low-mass galaxies are younger than high-mass galaxies. Previous studies of the rate of evolution of field early-type galaxies found a large range of mutually exclusive values. We show that these differences are largely caused by the differences between fitting methods: most literature studies are consistent with our result and with one another when using the same method. Finally, five of the early-type galaxies in our sample have AGNs. There is tentative evidence that the stellar populations in these galaxies are younger than those of galaxies without AGNs.

Subject headings: cosmology: observations — galaxies: evolution — galaxies: formation — galaxies: fundamental parameters — galaxies: kinematics and dynamics

Online material: color figures

1. INTRODUCTION

Understanding the formation and evolution of early-type galaxies is a key issue when addressing the mass assembly and star formation history of the galaxy population as a whole and the formation of structure in the universe, as 50% or more of all stars in the present-day universe are in early-type galaxies and bulges (see, e.g., Bell et al. 2003).

In hierarchical galaxy formation theories (e.g., Cole et al. 2000), massive galaxies assemble late, such that strong evolution of the mass density from $z = 1$ to the present day is expected (see, e.g., Kauffmann & Charlot 1998). Measuring the mass density requires a measurement of the luminosity density and an accurate determination of the $M/L$. $M/L$ can be estimated from models (see, e.g., Bell et al. 2004), but these estimates are uncertain due to the age/metallicity degeneracy and the unknown initial mass function (IMF) of the stellar populations of the galaxies (Bruzual & Charlot 2003).

The fundamental plane (FP; Djorgovski & Davis 1987; Dressler et al. 1987) provides a tool to measure the evolution of $M/L$ without model uncertainties. The $M/L$ offset of high-redshift galaxies from the local FP can be used to calibrate
high-redshift galaxy masses and to estimate the age of their stellar populations (Franx 1993). This technique has been used successfully to measure the luminosity-weighted ages of massive cluster galaxies, which have formed most of their stars at redshifts \( z \geq 2 \) (see, e.g., van Dokkum & Franx 1996; van Dokkum & Stanford 2003; Holden et al. 2005). However, it is not clear whether galaxies in the general field evolve in the same way. In fact, the hierarchical picture the formation redshift of galaxies with a given mass depends on environment (Diaferio et al. 2001). This would lead to substantial age differences between field and cluster galaxies at any redshift (van Dokkum et al. 2001). Since this is a generic property of all hierarchical formation models, measuring this difference is a critical test for those theories.

Various authors have measured the \( M/L \) evolution of field early-type galaxies through deep spectroscopy of magnitude-limited samples. The results are much less conclusive than the results from cluster studies, and the comparison between field and cluster has proved to be very hard. Some authors claim much faster evolution for field galaxies than for cluster galaxies (Treu et al. 2002; Gebhardt et al. 2003, hereafter G03), but others find that field and cluster galaxies evolve at comparable rates (van Dokkum et al. 2001; van Dokkum & Ellis 2003; van der Wel et al. 2004). Studies involving lensing galaxies (Kochanek et al. 2000; Rusin et al. 2003; van de Ven et al. 2003) indicate the presence of a mix of fast and slowly evolving galaxies. It is unclear whether the differences between the various results are caused by selection effects, measurement errors due to low signal-to-noise ratio (S/N) spectra, low number statistics, or contamination by late-type galaxies.

This paper describes a study of early-type galaxies at \( z \approx 1 \) using much higher quality data than in previous studies. The substantially large number of objects with very high S/N spectra enables us to accurately measure the \( M/L \) evolution of the field early-type galaxy population, to compare the cluster and field populations, to study correlations between \( M/L, M, \) and rest-frame color, and to describe the possible effects of biases. Also, we carefully compare the samples and results from previous studies and this study in order to verify previous claims about the evolution of field galaxies and to see whether previous results are in fact consistent with each other and these new results.

In § 2 we describe the sample selection, the spectroscopic observations and data reduction, and the measurement of velocity dispersions. In § 3 we describe the measurement of the structural parameters, colors, morphologies, and the available X-ray data. In § 4 we present the results. Throughout this paper we use Vega magnitudes and assume \( (\Omega_m, \Omega_\Lambda) = (0.3, 0.7) \), with a Hubble constant of \( H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1} \).

2. SPECTROSCOPY

2.1. Sample Selection and Observations

We selected galaxies in the Chandra Deep Field–South (CDFS) and the RDCS 1252.9–2927 cluster field (CL 1252; Rosati et al. 2004), which both have deep optical imaging from the Advanced Camera for Surveys (ACS) on the Hubble Space Telescope (HST). GOODS\(^1\) provides publicly available imaging in four filters (Giavalisco et al. 2004): F475W, F606W, F775W, and F850LP (hereafter \( b, v, i, \) and \( z \)). As these data were not yet available when we started this project, we used ground-based COMBO-17 photometry (Wolff et al. 2004) to select our sample for the first observing run. For subsequent runs version 0.5 of the ACS GOODS data was available, and for the last run we used the version 1.0 data release. Blakeslee et al. (2003) provide ACS imaging on the CL 1252 field in the \( i \) and \( z \) bands.

In order to construct a sample of early-type galaxies at \( z \approx 1 \) in the CDFS, we selected objects with \( i - z > 0.86 \) and COMBO-17 photometric redshifts in the range \( 0.8 < z_{\text{phot}} < 1.4 \). (We use \( z \) when we mean redshift and \( z_{\text{mag}} \) if we mean \( z \)-band magnitude, but when indicating a color we omit the mag subscript for clarity.) This color cut selects galaxies redder than a local Sbc galaxy at \( z = 1 \). Therefore, this study only includes galaxies that are on the red sequence at \( z \approx 1 \). We morphologically classified all galaxies satisfying these criteria and brighter than \( z_{\text{mag}} = 21.5 \), distinguishing between early- and late-type galaxies using the ACS imaging. The classification was based on compactness, regularity, and the presence of spiral arms. We classified 26 out of the 52 galaxies satisfying our selection criteria as early-type galaxies. We designed multi-slit masks for three different pointings, selected by the number of primary targets that could be included. Open spaces in the masks were filled with early-type galaxies satisfying the color and redshift criteria but fainter than \( z_{\text{mag}} = 21.5 \), early-type galaxies with lower photometric redshifts, and late-type galaxies with \( i - z > 0.85 \) and redshifts \( 0.8 < z_{\text{phot}} < 1.4 \). Switching from using ground-based \( i - z \) colors from COMBO-17 to ACS \( i - z \) colors from GOODS did not lead to large differences between the selected samples, although several objects changed priority.

The same selection criteria were used for the CL 1252 field. However, the mask design for the CL 1252 field was geometrically constrained because the primary targets were cluster galaxies at \( z = 1.24 \). The two brightest galaxies in the cluster are the two central galaxies, which had to be included in a single slit because of their small angular separation of \( 1.5' \). Therefore, not only the positions but also the position angles of the designed masks were fixed. Unfortunately, only two galaxies brighter than \( z_{\text{mag}} = 21.5 \), redder than \( i - z = 0.85 \), and with early-type morphologies could be included, additionally to the cluster galaxies. Similarly to the CDFS masks, fillers were included.

We carried out the observations with FORS2 in MXU mode on ESO’s Kueyen, one of the VLT unit telescopes. We used the 600z grism together with the OG590 order separation filter to obtain a sufficiently high spectral resolution (\( \sigma = 80 \text{ km s}^{-1} \)) and to cover the wavelength range around the Balmer/4000 Å break for galaxies at \( z \approx 1 \). The observations were carried out in series of four dithered exposures with spatial offsets of \( 1.5'' \) or \( 2'' \) and equal exposure times ranging from 14 to 30 minutes each.

In total, 51 hr of scientifically useful integration time was acquired, of which 38 hr had seeing better than 1'' (\( \sigma = 80 \text{ km s}^{-1} \)). The cumulative integration time for the three pointings in the CDFS is 27 hr, with a median seeing of 0.95. The single pointing in the CL 1252 field has an integration time of 24 hr, with a median seeing of 0.65. These observations were carried out during five different observing runs from 2002 September to 2003 November.

The sample described in this paper consists of 38 galaxies with velocity dispersions, of which 20 are early-type galaxies at \( z \approx 1 \) and 18 are early-type galaxies at lower redshift, or late-type galaxies. One hundred percent of our primary targets yielded velocity dispersions. The CL 1252 observations also yielded four velocity dispersions of cluster galaxies. The FP of the CL 1252 cluster is discussed by Holden et al. (2005).

2.2. Data Reduction

The spectroscopic data were reduced using standard IRAF tasks. Lamp flat fields were taken before or after each night, in
sequences of five exposures. We used the sequence closest in time to the science observation. Cosmic rays were removed using the L.A.Cosmic task (van Dokkum 2001). Afterward, all frames were checked manually. We subtracted a two-dimensional sky spectrum from each exposure, obtained by median averaging the four dithered exposures in a sequence, masking the target and secondary or serendipitous objects, if present. The atmospheric emission lines, which are bright and abundant in the observed wavelength range, were used to perform the wavelength calibration. We corrected for distortion in the spatial direction by tracing the target. All individual exposures were optimally weighted to obtain maximum S/N.

There are various atmospheric absorption features in the observed wavelength range. Because the strength and shape of these features change with air mass and atmospheric conditions, we needed to correct each exposure separately. To this end, we included a blue star in each of our masks, which was reduced along with the galaxy spectra. After the final combination, the regions in the galaxy spectra with atmospheric features were divided by the normalized spectrum of the blue star. Spectroscopic standard stars were used to do a relative flux calibration. One-dimensional spectra were extracted by adding those pixel rows with more than 25% of the flux of the brightest row, weighting optimally.

The smoothed one-dimensional spectra are shown in Figures 1 and 2. The coordinates of the objects for which we measured velocity dispersions (see § 2.3) are given in Table 1. Redshifts, S/N, and emission lines are given in Table 2.

2.3. Velocity Dispersions

Velocity dispersions are obtained by fitting template spectra to the observed galaxy spectra. The fitting method is extensively described by van Dokkum & Franx (1996). The continua of both the observed and the template spectra are filtered out in
Fig. 1.—Continued

CDPS-3 (z=1.044)

CDPS-4 (z=0.964)

CDPS-5 (z=0.685)

CDPS-6 (z=0.660)

CDPS-7 (z=1.155)

CDPS-8 (z=1.125)

CDPS-9 (z=1.097)

CDPS-10 (z=1.119, e)

CDPS-11 (z=1.096, e)

CDPS-12 (z=0.955)

CDPS-13 (z=0.669)

CDPS-14 (z=1.096)
Fig. 1.—Continued
Fourier space and the template spectrum is convolved with a Gaussian to match the width of the absorption lines in the galaxy spectrum. The part of the galaxy spectrum used in the fit is as large as possible. Therefore, our measurements do not rely on a few high-S/N absorption features.

As templates we use coude spectra of 132 stars with the appropriate wavelength range from the sample constructed by Valdes et al. (2004), with a spectral range from F0 to M6, including both very low and high metallicity stars, and different luminosity classes. These spectra have an FWHM resolution of about 1 Å. Each stellar spectrum needs to be smoothed to each galaxy spectrum separately before being rebinned. A second-order function is fitted to the width of atmospheric emission lines as a function of wavelength to obtain the spectral resolution to which the template spectra are smoothed.

When fitting the galaxy spectra, we weight with the inverse of the sky brightness, and we mask the region around the atmospheric A band at 7600 Å. The spectrum above 9300 Å is omitted because of the strong atmospheric absorption, the ever increasing brightness of the sky emission lines, and the decreasing system throughput.

After performing the fit for a small number of templates, masking, and weighting as described, we check the residuals from the fit. Regions with emission lines, large sky line residuals, and remaining data artifacts such as cosmic-ray remnants are masked if present. We then fit the galaxy spectrum with all template spectra. We check whether the obtained parameters change strongly if one or two strong features are masked out, but we conclude that this generally is not the case: excluding the strongest features from the fit increases the $\chi^2$ value but does not change the results significantly in most cases. For some spectra, however, including Balmer lines in the fit leads to different results, probably because unseen emission-line contributions contaminate these features. For low-quality spectra (S/N $\leq$ 10 per 1.6 Å pixel in the extracted, one-dimensional spectra) the contributions of Balmer lines or other strong features can hardly be checked because excluding these strong features leaves insufficient signal to obtain a proper fit. Therefore, we exclude objects with S/N < 12 spectra from our analysis, but we mention the effect of including these.

For all spectra with S/N $\geq$ 12 the random errors are below 3% for $\sigma > 200$ km s$^{-1}$ and below 5% for $\sigma < 200$ km s$^{-1}$. Adding a systematic uncertainty (including template mismatch and the error on the resolution of the galaxy spectra) of about 10% for S/N = 10 spectra and 2% for the highest S/N spectra, the total errors range from 3% to 17% with a median of 7.5% for our sample of early-type galaxies with S/N $\geq$ 12. Some galaxies have measured velocity dispersions that are not much larger than the resolution of the spectra. Although these are included in the analysis, they play no important role in the derivation of our results.

The best-fitting stellar spectral type and the measured velocity dispersion are listed in Table 1. These velocity dispersions are aperture corrected to a 3.4 diameter circular aperture at the distance of Coma as described by Jørgensen et al. (1995). This correction ranges from 5% to 7%.

3. PHOTOMETRY

3.1. Profile Fitting and Morphologies

The ACS provides us with an unprecedented combination of deep and high-resolution imaging. The spatial resolution (FWHM) at $z = 1$ is 0.8 kpc, allowing us to accurately measure the effective radii of early-type galaxies at this redshift, which typically are a few kiloparsecs. We use the single, unstacked, flat-fielded frames publicly available through the HST MAST.
For the CDFS the number of frames for different positions ranges from 8 to 24 (with 530 s exposure time each). For the CL 1252 field the number of frames ranges from 10 to 40 (with 1200 s exposure time each) but is mostly 10 as only the center of the cluster has 40 overlapping images.

For each galaxy each individual z-band image is fitted by a 1/n model (with n = 1, 2, 3, 4) convolved by a position-dependent point-spread function created with Tiny Tim (Krist 1995), measuring r_eff (the effective radius), \( \mu_{\text{eff}} \) (the surface brightness at \( r_{\text{eff}} \)), the position angle, and the ellipticity. Each individually derived set of model parameters is distortion corrected by calculating the pixel scales in the x- and y-directions, using the polynomial distortion coefficients available online.\(^2\) We then average the results and compute the measurement error from the scatter. The error is generally about 6% in \( r_{\text{eff}} \), but the combination of the uncertainty in \( r_{\text{eff}} \) and \( \mu_{\text{eff}} \) is such that it is directed almost parallel to the local FP. The error relevant to the offset from the FP is typically 2%. Thus, uncertainties in the offset from the local FP are dominated by the uncertainty in \( \sigma \), of which the error is pointed almost perpendicular to the FP. The effective radii and surface brightnesses are given in Table 1. For consistency with earlier studies, these are the values obtained from fitting a de Vaucouleurs profile in our analysis.

To transform the observed z-band surface brightnesses to the rest-frame \( B \) band, we use the technique described by van Dokkum & Franx (1996), using the templates from Coleman et al. (1980) and observed colors (see \( \S \) 3.2) to interpolate between the passbands. The calculated rest-frame \( B \)-band surface brightnesses only depend very weakly on the spectral type of the template used. The typical difference found for using the Sbc template instead of the E template is less than 0.02 mag. The transformations are a function of redshift. As an example, we give the transformation (based on the E template) for a galaxy at \( z = 1 \):

\[
B_z = z + 0.165(i - z) + 1.398. \tag{1}
\]

Physical sizes and rest-frame \( B \)-band surface brightnesses are given in Table 3.

Figures 3 and 4 show the combined residuals of the \( r^{1/4} \) fits along with the color images of the 38 galaxies with velocity dispersions. Figure 3 shows the early-type galaxies, Figure 4 the late-type galaxies. Two numbers are used to characterize the magnitude of the residuals. At the upper right of each residual image the absolute value of the flux in the asymmetric part of the residual is given as a percentage of the total flux of the galaxy, obtained by subtracting the residual rotated by 180° from the residual itself. In the upper left the absolute value of the flux in the symmetric part of the residual is given as a percentage of the total flux of the galaxy. We find that the asymmetric residual is a good indicator of morphology. Our final morphological qualification is a combination of the magnitude of the asymmetric residual and visual inspection of the cause of the asymmetry. Some of the early-type galaxies (Fig. 3) have significant asymmetric residuals, but these are caused by features on a very small scale, in the centers of the galaxies, i.e., not by features attributed to spiral arms or other large-scale irregularities. The late-type galaxies (Fig. 4) have large asymmetric residuals, caused by large-scale structures like spiral arms. We note that the Sercial number does not distinguish well between late- and early-type galaxies. For example, the most massive galaxy in our sample has \( n = 2 \), and some galaxies that we classify as late-type galaxies have \( n = 4 \).

### 3.2. Colors

We supplement the ACS imaging with ground-based optical and near-IR imaging from FORS2 and ISAAC on the Very Large Telescope (VLT) and SOFI on the New Technology Telescope (NTT; B. Vandame et al. 2005, in preparation). GOODS ACS imaging of the CDFS provides photometry in the \( b \), \( v \), \( i \), and \( z \) bands (data release ver. 1.0), and ESO’s imaging survey\(^3\) provides SOFI and ISAAC imaging in the \( J \) and \( K \) bands. Since the CDFS is not entirely covered by ISAAC, we use the SOFI data for the objects outside the ISAAC pointings. All images were smoothed to match the resolution of the \( K \)-band data with the worst seeing, which is 0.8′ for the ISAAC imaging and 1′′ for the SOFI imaging. The photometric differences between the ISAAC and SOFI data sets are small (<0.01 mag), since the zero points of the ISAAC data are based on SOFI photometry.

\(^{2}\) See http://www.stsci.edu/hst/acs/analysis/PAMS.

\(^{3}\) See http://www.eso.org/science/eis.
the fifth column is per CDFS-29 .............. CDFS-28

CDFS-24 .............. CDFS-27 .............. O

CDFS-21 .............. CDFS-16 .............. CDFS-15 .............. CDFS-14 .............. CDFS-5 ..............

CL 1252-8............ O

colors. To transform the observed colors to rest-frame not only

For both the CDFS and the CL 1252 field deep X-ray data from Chandra are available, such that we can check for the presence of active galactic nuclei (AGNs) in our galaxy sample. Gioacch et al. (2002) and Alexander et al. (2003) provide catalogs of the CDFS data. The Chandra data of the CL 1252 field are described by Rosati et al. (2004), who also constructed a point-source catalog. Eight galaxies in our sample of 38 are
identified as AGNs, based on their large X-ray luminosities (typically $>10^{42}$ ergs s$^{-1}$). Five of these are early-type galaxies, of which two have emission lines in their spectra. Besides the eight AGNs, two galaxies in our sample are identified as extended X-ray sources. This X-ray radiation is accounted for by diffuse halo gas. The X-ray luminosities of CDFS-4 and CDFS-22 are $7.15 \times 10^{41}$ and $3.42 \times 10^{42}$ ergs s$^{-1}$, respectively. It is not surprising that CDFS-4 and CDFS-22 turn out to be two of the most massive galaxies in our sample. Also, CDFS-22 is one of the galaxies with an AGN. The X-ray properties of our sample of most massive galaxies in our sample. Also, CDFS-22 is one of the most massive galaxies in our sample. Also, CDFS-22 is one of the most massive galaxies in our sample. Also, CDFS-22 is one of the most massive galaxies in our sample. Also, CDFS-22 is one of the most massive galaxies in our sample. Also, CDFS-22 is one of the most massive galaxies in our sample. Also, CDFS-22 is one of the most massive galaxies in our sample. Also, CDFS-22 is one of the most massive galaxies in our sample. Also, CDFS-22 is one of the most massive galaxies in our sample. Also, CDFS-22 is one of the most massive galaxies in our sample. Also, CDFS-22 is one of the most massive galaxies in our sample. Also, CDFS-22 is one of the most massive galaxies in our sample. Also, CDFS-22 is one of the most massive galaxies in our sample. Also, CDFS-22 is one of the most mass...
derive effective radii separately for each band, we compute the surface brightness in the $r$ band at the effective radius as measured in the $g$ band. We note that the FP coefficients as derived by Bernardi et al. (2003) are different from those from Jørgensen et al. (1996), but this does not lead to different results. The reason we use the FP coefficients as derived from the Faber et al. (1989) data, and not the Bernardi et al. (2003) data, is that Faber et al. (1989) use $B$-band photometric data.

In Figure 5 we show the FP of the SDSS galaxies and our sample, where the surface brightnesses of all galaxies are transformed to the value they would have at $z = 1$, assuming luminosity evolution found for massive cluster galaxies, $\Delta \ln (M/L_B) = -1.12z$. This value for the $M/L$ evolution of massive cluster galaxies is derived from compiling all existing data in the literature for galaxies more massive than $M = 2 \times 10^{11} M_\odot$ (van Dokkum & Franx 1996; Kelson et al. 2000; van Dokkum & Stanford 2003; Wuyts et al. 2004; Holden et al. 2005). Our field sample shown in Figure 5 includes all early-type galaxies with spectra with $S/N \geq 12$. This is also the sample used in the analysis throughout the rest of the paper and in the subsequent figures. As can be seen, the FP already existed at $z = 1$ for a large range in size. At low masses outliers occur, but the interpretation is not straightforward, as selection effects play a major role in this regime (see § 4.3).

4.2. Evolution of $M/L$ with Redshift

The offset of high-redshift galaxies from the local FP is interpreted as a difference in $M/L$ as compared to equally massive
Fig. 3.—Continued

Fig. 4.—Color images and s1/4 profile fit residuals of the galaxies with late-type or irregular morphologies. These galaxies have measured velocity dispersions but are not included in the analysis in the subsequent sections. For an explanation of the images and the numbers, see Fig. 3. [See the electronic edition of the Journal for a color version of this figure.]
local galaxies (e.g., van Dokkum & Franx 1996). Figure 6 shows the offsets of our field galaxy sample in \(\Delta \ln (M/L_B)\) as a function of redshift. The values are listed in Table 3. Cluster samples from the literature are also shown. The galaxies in our sample seem to evolve faster than the galaxies in the cluster samples, and the scatter in \(\Delta \ln (M/L_B)\) is large. Before we interpret the scatter and the apparent difference between field and cluster galaxies, we need to investigate the origin of the scatter. In Figure 7 we show \(M/L_B\) as a function of the rest-frame \(I - B\) color. Galaxies with low \(M/L\) are bluer than galaxies with high \(M/L\), as expected from the stellar population models shown in the figure. In Figure 8 we show a similar relation between rest-frame \(U - B\) and \(M/L_B\), but in this case the correlation is less clear, due to the fact that the range of \(U - B\) colors is much smaller than the range of \(B - I\) colors, and probably also because \(U - B\) is more sensitive to small variations in the star formation history. Considering the correlation between color and \(M/L\) and the fact that \(M/L\) evolves with redshift, one expects that color evolves with redshift as well. In Figure 9 we show \(B - I\) as a function of redshift; note the strong similarities between Figures 6 and 9. The strong correlation between color and \(M/L\) and the similarity in \(M/L\) evolution and color evolution confirm that the observed evolution and scatter of \(M/L\) are intrinsic and not due to measurement errors.

We calculate the evolution of \(M/L\) of our field sample by performing a linear fit and minimizing the mean deviation, weighting by the inverse of the error, and forcing the fit to go through the \(z = 0.02\) data point derived from the Faber et al. (1989) sample. We separately consider the evolution of the primary sample, which contains the galaxies satisfying all of our selection criteria. We find that the average evolution of our entire early-type galaxy sample is \(\Delta \ln (M/L_B) = (-1.75 \pm 0.16)z\) \([-1.72 \pm 0.15\) for the primary sample alone], which is significantly faster than the evolution found for cluster galaxies, which is \(\Delta \ln (M/L_B) = (-1.28 \pm 0.08)z\) (van Dokkum & Franx 1996; Kelson et al. 2000; van Dokkum & Stanford 2003; Wuyts et al. 2004; Holden et al. 2005). The scatter in \(\Delta \ln (M/L_B)\) is 0.58 for our field galaxy sample (0.54 for the primary sample) and 0.28 for the MS 1054 cluster sample (Wuyts et al. 2004).

Figure 5 suggests that the \(M/L\) evolution may depend on galaxy mass, as galaxies with small \(r_{eff}\) and low \(\sigma\) tend to lie lower with respect to the local FP as compared to galaxies with large \(r_{eff}\) and high \(\sigma\). This was also found for cluster galaxies by Wuyts et al. (2004). We estimate \(M\) and \(M/L\) in solar units as described by van Dokkum & Stanford (2003). The values are listed in Table 3. There is a striking difference between low- and high-mass galaxies. For galaxies with masses \(M > 2 \times 10^{11} M_\odot\), we measure \(\Delta \ln (M/L_B) = (-1.20 \pm 0.18)z\) for our field sample and \((-1.12 \pm 0.06)z\) for the cluster samples. For the massive galaxies in the primary sample alone we find \(\Delta \ln (M/L_B) = (-1.26 \pm 0.18)z\). The observed scatter is decreased to 0.34 for the field sample (0.32 for the primary sample) and to 0.28 for the cluster samples. When changing this mass cut to \(3 \times 10^{11} M_\odot\), as is done by Wuyts et al. (2004), but thereby limiting the number of galaxies in our sample to 4, we find \((-1.12 \pm 0.13)z\) and \((-0.99 \pm 0.10)z\) for our sample and the cluster samples,
Fig. 7.—Rest-frame $B - I$ color vs. $M/L_B$ for our early-type galaxy sample. For an explanation of the symbols, see Fig. 6a. The large filled circle at the upper right indicates the median $B - I$ and $M/L_B$ of the massive galaxies ($M > 2 \times 10^{11} M_\odot$) in the SDSS field early-type galaxy sample. The orientation of the distribution around the median values and the amount of scatter are indicated by the tilted error bars. The dotted line is a solar metallicity Bruzual-Charlot model for a single stellar population. The solid line is a model with exponentially declining star formation ($\tau = 1$ Gyr). Both model tracks are shifted vertically to match the SDSS data point. Model ages are indicated by ticks at intervals of 1 Gyr. The correlation between $M/L$ and color implies that the observed scatter in $M/L$ is real and can be ascribed to age differences between the stellar populations of the galaxies. Our $i - z \geq 0.86$ color selection limit roughly corresponds to $B - I \geq 1.1$ according to the Bruzual-Charlot models. This shows that our selection criterion only excludes galaxies with ages less than 1 Gyr and does not affect our conclusions regarding the massive, red galaxies. [See the electronic edition of the Journal for a color version of this figure.]

respectively. We conclude that for high-mass galaxies, there is no difference between the cluster samples and our field sample. The galaxies with masses $M < 2 \times 10^{11} M_\odot$ in our sample evolve much faster: $\Delta \ln (M/L_B) = (-1.97 \pm 0.16) z \ [(1.90 \pm 0.17) z$ for the primary sample alone]. We verify that these results do not change if galaxies with spectra with $S/N < 12$ are included as well. Therefore, the accuracy of our results is not limited by the quality of the velocity dispersions.

In Figures 6 and 9 we show evolutionary tracks for a single stellar population for formation redshifts $z = 1$ and 2. This very simple model assumes that luminosity evolves with time as $L \propto (t - t_{\text{form}})^\kappa$, where $\kappa$ is derived from stellar population models (for more details see van Dokkum et al. 1998). These tracks indicate a large spread in formation redshifts. We note, however, that $\kappa$ is sensitive to the IMF. Here we use a single stellar population from Bruzual & Charlot (2003) with a Salpeter IMF and solar metallicity, which yields $\kappa = 0.97$ in the $B$ band and 0.43 in the $I$ band. According to this model, the massive galaxies have high luminosity-weighted formation redshifts ($z \geq 2$), whereas less massive galaxies in our sample have lower formation redshifts ($1 < z < 2$).

4.3. The Relation between $M$ and $M/L$ and the Role of Selection Effects

Figure 10 illustrates the tight relation between mass and $M/L_B$. We also show the SDSS field sample described in § 4.1.

Fig. 8.—Rest-frame $U - B$ color vs. $M/L_B$ for the early-type galaxy sample. For an explanation of the symbols, see Fig. 6a. The “LOCAL” data point is taken from G03. There is a similar relation as in Fig. 7, but it is less clear because the range of colors is much smaller in $U - B$ than in $B - I$. Our $i - z \geq 0.86$ color criterion corresponds to $U - B \geq 0.07$ at $z = 1$, which demonstrates that this criterion would only exclude the most extremely blue galaxies (with ages well below 1 Gyr). [See the electronic edition of the Journal for a color version of this figure.]

Fig. 9.—Evolution of the rest-frame $B - I$ color with redshift. For an explanation of the symbols, see Fig. 6a. As already suggested by the tight relation between $B - I$ and $M/L$ in Fig. 7, the evolution of color with redshift is very similar to the evolution of $M/L$ with redshift. Massive galaxies are the reddest in the local universe, and their color evolves slower than the color of low-mass galaxies. Also, the color of the most massive field galaxies is very similar to the color of massive cluster galaxies. [See the electronic edition of the Journal for a color version of this figure.]
M and M/LB are calculated as described by van Dokkum & Stanford (2003): log M = 2 log σ + log r_{eff} + 6.07, log (M/LB) = 2 log σ + 0.4(r_{eff,B} - 27) - log r_{eff} - 1.29. The M/L of all galaxies have been corrected for evolution as derived from massive cluster galaxies, Δ ln (M/LB) = -1.12z.

At z = 1, the relation between M/L and M seems much steeper than in the local universe. However, we have to take selection effects into account. Since we selected our sample in the z band, we can transform our magnitude limit into a luminosity limit at z = 1. The dotted line indicates our magnitude limit, translated into a luminosity limit at z = 1. Therefore, this limit only applies to the filled circles. The solid line indicates the median M/LB of the SDSS early-type galaxy population. The long-dashed line indicates the median M/LB of galaxies that are brighter than our luminosity limit at z = 1. Assuming that the scatter in M/LB at z = 1 is a factor of 2 larger than in the local universe, the median of the M/L of galaxies brighter than the luminosity limit follows the short-dashed line. It is clear that for the three galaxies in the primary sample with the lowest masses selection effects play such a dominant role that we cannot include these objects in our efforts to correct for this bias. Between 6 × 10^{10} and 2 × 10^{11} M_\odot, selection effects are relevant but not dominant. For higher masses, selection effects do not affect our sample. The difference between galaxies with masses of ~10^{11} M_\odot and those with masses of ~10^{12} M_\odot cannot be explained without an increase in the scatter with redshift or assuming mass-dependent ages. [See the electronic edition of the Journal for a color version of this figure.]

Fig. 10.—M vs. M/LB as derived from the FP for the early-type galaxies in our sample and for the nearby sample from the SDSS. For an explanation of the symbols, see Fig. 6a. All data points have been corrected for M/L evolution as found for massive cluster galaxies, Δ ln (M/LB) = -1.12z, normalizing at z = 1. The dotted line indicates our magnitude limit, translated into a luminosity limit at z = 1. Therefore, this limit only applies to the filled circles. The solid line indicates the median M/LB of the SDSS early-type galaxy population. The long-dashed line indicates the median M/LB of galaxies that are brighter than our luminosity limit at z = 1. Assuming that the scatter in M/LB at z = 1 is a factor of 2 larger than in the local universe, the median of the M/L of galaxies brighter than the luminosity limit follows the short-dashed line. It is clear that for the three galaxies in the primary sample with the lowest masses selection effects play such a dominant role that we cannot include these objects in our efforts to correct for this bias. Between 6 × 10^{10} and 2 × 10^{11} M_\odot, selection effects are relevant but not dominant. For higher masses, selection effects do not affect our sample. The difference between galaxies with masses of ~10^{11} M_\odot and those with masses of ~10^{12} M_\odot cannot be explained without an increase in the scatter with redshift or assuming mass-dependent ages. [See the electronic edition of the Journal for a color version of this figure.]

M and M/LB are calculated as described by van Dokkum & Stanford (2003): log M = 2 log σ + log r_{eff} + 6.07, log (M/LB) = 2 log σ + 0.4(r_{eff,B} - 27) - log r_{eff} - 1.29. The M/L of all galaxies have been corrected for evolution as derived from massive cluster galaxies, Δ ln (M/LB) = -1.12z.

At z = 1, the relation between M/L and M seems much steeper than in the local universe. However, we have to take selection effects into account. Since we selected our sample in the z band, we can transform our magnitude limit into a luminosity limit in the rest-frame B band, which is close to the observed z band at z ~ 1 (the median redshift of our sample of 16 z ~ 1 galaxies is z = 1.04). Because this luminosity limit only applies to our primary sample, this discussion does not involve the fillers (mainly galaxies at z ~ 0.7). We comment further on these galaxies below.

We first test whether the observed distribution can be fully explained by selection effects, assuming that the slope and the scatter of the FP do not evolve. The probability that the observed distribution is drawn from a population with the same distribution as the SDSS galaxies is 0.14%, if the M/L evolution is the same for all galaxies. Therefore, the slope evolves, the scatter evolves, or both. If there is an age difference between high- and low-mass galaxies, the scatter will most probably also evolve differently for high- and low-mass galaxies. We cannot exclude with high confidence that the observed distribution of M/L is due to a larger scatter at high redshift: if the scatter in M/L at z = 1 is twice as large as in the local universe, the probability of a non-evolving slope is 8.0%. However, we see no evidence for an increase in the scatter at the high-mass end, where selection effects do not play a role. Hence, if the scatter evolves, this is only true for galaxies with masses M ~ 10^{11} M_\odot. Since the increased scatter is most likely caused by young ages, low-mass galaxies would have lower M/L, and hence the slope of the relation would also be changed. We consider these findings as strong evidence for mass-dependent evolution of early-type galaxies. To confirm that we observed a change in the slope of the FP, deeper and larger surveys are needed.

The location of the luminosity limit at z = 1 in Figure 10 shows that our z ~ 1 sample is dominated by selection effects for masses M < 6 × 10^{10} M_\odot. The objects with such low masses are only included in the sample because of their probably extreme M/L. We cannot correct the M/L of that subsample for the bias introduced by our luminosity limit.

On the other hand, for the sample of galaxies with higher masses, we can correct for selection effects because they are relevant, but not dominant, as can be seen in Figure 10. The average evolution of the galaxies in the primary sample with masses M > 6 × 10^{10} M_\odot is Δ ln (M/LB) = (-1.55 ± 0.16)z. The median mass of this subsample of 12 galaxies is M = 1.9 × 10^{11} M_\odot. We estimate the maximum bias by assuming that the slope is the same at z = 1 and in the local universe, but that the scatter is a factor of 2 larger at z = 1 than at z = 0. In that case the observed distribution is expected to follow the short-dashed line in Figure 10. At a given mass, the difference between the solid line and the short-dashed line is the bias introduced by the selection effects in luminosity. We increase the observed M/L of each galaxy in our primary sample by the difference between the solid line and the short-dashed line at the mass of that galaxy. For galaxies more massive than M = 2 × 10^{11} M_\odot, this correction is negligible, but for the galaxies with masses M ≈ 6 × 10^{10} M_\odot this correction is about 30%. Using this method, we find a bias-corrected evolution of the galaxies with masses M > 6 × 10^{10} M_\odot of Δ ln (M/LB) = (-1.43 ± 0.16)z. Given the uncertainty in the intrinsic scatter, deeper observations are necessary to confirm this value.

Besides a bias due to the luminosity limit, errors in the velocity dispersion produce correlated errors in M and M/L, hence understimating the measured evolution. Taking the errors in σ into account, we find that this introduces a bias at the level of only 2%–3%, which is several times smaller than our measurement accuracy.

The above analysis only involves z ~ 1 galaxies satisfying all our selection criteria, but we note that the relation between M/L and M exists for the z ~ 0.7 galaxies as well. However, this subsample is selected in an inhomogeneous way; therefore, it is impossible to correct the observed evolution. Since all the galaxies roughly lie along lines of constant luminosity, the bias toward low-M/L galaxies likely explains the observed relation between M and M/L for this subsample.

Besides by luminosity and morphology, the galaxies in our sample are also selected by color. This potentially introduces an important bias in the measured evolution because of the exclusion of galaxies with blue colors, i.e., low M/L. Our color criterion, however, is quite generous. Even the very blue, low-mass galaxies satisfy this criterion. For typical Bruzual-Charlot models, our color limit (i - z = 0.86) corresponds to U - V ~ 0.67 at z = 1, which is 0.2 mag bluer than the limit applied by Bell et al. (2004) to select galaxies on the red sequence at
**Comparison with Previous Results**

| REFERENCES | Reported | Fitted, High Mass | Fitted, Low Mass | (log (M/L_B)) | (z) | N  |
|------------|----------|------------------|-----------------|--------------|-----|----|
| Treu et al. (2001) | -1.64 ± 0.12 | -1.64 ± 0.34 | -1.55 ± 0.27 | -1.47 ± 0.89 | 2.3 × 10^{11} | 0.29 | 19 |
| van Dokkum et al. (2001) | -1.35 ± 0.35 | -1.67 ± 0.23 | -1.29 ± 0.39 | -1.94 ± 0.33 | 1.5 × 10^{11} | 0.42 | 18 |
| Treu et al. (2002) | -1.84 | -1.25 ± 0.25 | -1.68 ± 0.13 | -1.41 ± 0.29 | -1.77 ± 0.24 | 1.3 × 10^{11} | 0.56 | 27 |
| van Dokkum & Ellis (2003) | -1.25 ± 0.25 | -1.19 ± 0.20 | -1.15 ± 0.49 | -2.05 ± 0.17 | 8.5 × 10^{10} | 0.64 | 21 |
| G03 | -0.30 | -0.22 | -0.31 | -0.23 | 1.6 × 10^{11} | 0.90 | 27 |
| This paper | -1.75 ± 0.16 | -1.75 ± 0.16 | -1.20 ± 0.18 | -1.97 ± 0.16 | 1.6 × 10^{11} | 0.90 | 27 |
| Rusin et al. (2003) | -1.24 ± 0.21 | ... | ... | ... | ... | 0.54 | 21 |
| van de Ven et al. (2003) | -1.43 ± 0.30 | -1.36 ± 0.18 | -1.13 ± 0.31 | -1.71 ± 0.30 | 2.3 × 10^{11} | 0.54 | 21 |

**Notes.**—Col. (2): Values of the evolution of M/L as reported in the cited papers. Col. (3): Values we derive using our fitting technique on the tabulated data of the individual galaxies in the cited papers. Col. (4): Evolution of galaxies more massive than $M = 2 \times 10^{11} M_{\odot}$, using our fitting technique. Col. (5): Evolution of less massive galaxies. Average masses, redshifts, and sample sizes ($N$) of the samples are also listed. When two references are given, the results presented in the newest paper are based on data of both papers. Treu et al. (2002) only list redshifts and velocity dispersions of the galaxies. The reported values are all uncorrected for selection effects. Treu et al. (2002) only give an error for the uncorrected value, and G03 do not give an error. Besides the magnitude-limited samples, the two studies of the same sample of lensing galaxies are also listed. We only fitted the tabulated data from van de Ven et al. (2003), but the data presented by Rusin et al. (2003) yield the same results.

4.5. Ongoing Star Formation and Morphological Deviations

It is interesting to see that galaxies with emission lines have relatively low $M/L$, as can be seen from both Figures 6 and 10. Treu et al. (2002) have already shown that a substantial fraction of the massive early-type galaxy population at high redshift shows evidence for ongoing star formation by the presence of emission lines. Galaxies with emission lines in our sample, however, tend to have low masses. Figure 10 suggests that several of the galaxies with masses $\sim 3 \times 10^{10} M_{\odot}$ are included in the magnitude-limited sample only because they are forming stars. When excluding galaxies with emission lines, the evolution of the sample (without applying a mass cut) is $\Delta \ln (M/L_B) = (-1.61 \pm 0.18)z$. The galaxies with emission lines evolve faster: $\Delta \ln (M/L_B) = (-2.41 \pm 0.45)z$. On average, the mass of the galaxies without emission lines is 1.5 times larger than the mass of the galaxies with emission lines.

Mergers or interactions, accompanied by star formation, can lead to both deviations from smooth $p^{\text{hi}}$ profiles and low $M/L$ values. Van Dokkum & Ellis (2003) have shown tentative examples of this phenomenon. Figure 11 shows the magnitude of the asymmetric residual (described in \S 3.1), versus the offset from the FP, corrected for luminosity evolution. There is no correlation between deviations from smooth surface profiles and $M/L$, and we find no evidence for a connection between star formation activity and interactions or mergers.

4.6. AGNs

Four out of 11 early-type galaxies at $z < 0.8$ have AGNs, and 1 out of 16 for the $z > 0.9$ sample, as determined from X-ray imaging (see \S 3.3). Of the five early-type galaxies with AGNs, three are more massive than $M = 2 \times 10^{11} M_{\odot}$ and two have emission lines in their spectra. The $M/L$ evolution of galaxies in our sample with AGNs is $\Delta \ln (M/L_B) = (-1.64 \pm 0.31)z$, which is not very different from the value for galaxies without AGNs: $(-1.76 \pm 0.20)z$. However, the galaxies with AGNs are twice as massive as the galaxies without AGNs. Considering high-mass galaxies only ($M > 2 \times 10^{11} M_{\odot}$), we find $(-1.46 \pm 0.24)z$ for galaxies with AGNs and $(-1.02 \pm 0.41)z$ for galaxies without AGNs. Now, the trend is reversed, but only one out of the five galaxies with an AGN has a particularly blue

Our color cut is 0.45 mag bluer than the color-magnitude relation found by the same authors. From Bruzual-Charlot models we estimate that we only miss galaxies that are younger than $\sim 1$ Gyr (see Figs. 7 and 8). Furthermore, Bell et al. (2004) did not find blue, massive galaxies in the entire COMBO-17 data set. Hence, it is very unlikely that we miss any blue galaxy at the bright end of our sample because of our color selection criterion.

4.4. Independent Evidence for Mass-dependent Evolution of Early-Type Galaxies

It is particularly interesting to compare our results to the results from studies involving lensing galaxies, since those samples are mass selected. Rusin et al. (2003) and van de Ven et al. (2003) find $\Delta \ln (M/L_B) = (-1.29 \pm 0.09)z$ and $(-1.43 \pm 0.30)z$, respectively, using the same data set. This seems somewhat low compared to the evolution of our sample and results found in the literature, but the lensing galaxies typically have high masses (see Table 4). The data for individual galaxies published by van de Ven et al. (2003) show that the median mass of the lens sample is $M = 2 \times 10^{11} M_{\odot}$, whereas the median mass of our early-type galaxy sample is $M = 1.3 \times 10^{11} M_{\odot}$. The galaxies in the lens sample that are more massive than $M = 2 \times 10^{11} M_{\odot}$ evolve as $\Delta \ln (M/L_B) = (-1.13 \pm 0.31)z$. The galaxies less massive than this evolve much faster: $(-1.71 \pm 0.29)z$. The similarity between the results from the lensing sample and our sample is striking, especially because the lensing sample is mass selected. Hence, it is very hard to see how a bias toward low-$M/L$ galaxies can be responsible for the observed mass dependence in the lensing sample. It is possible that not all the low-mass lenses are genuine early types. If we omit the most irregular lenses (FBQ 0951+2635, SBS 1520+530, and B1608+656), we still find rapid evolution, $(-1.64 \pm 0.24)z$, for low-mass lensing galaxies. Hence, these results provide strong evidence that the observed dependence of $M/L$ evolution on mass in our sample is real. One possible complicating factor is that the lensing cross section of galaxies in groups is larger than that of galaxies in the lowest density environments. This may lead to a difference between the population of the lensing sample and the population of our sample.
to $10^{12} M_\odot$ and are evenly distributed in log $M$. Remarkably, when taking the tabulated data of the individual galaxies and applying our fitting method, we find a larger value: $(-1.67 \pm 0.23) z$ (see Table 4). Van Dokkum et al. (2001) create redshift bins in which they calculate the biweight center of the $M/L$ offset. This method is sensitive to the bin choice, which leads to the difference between the reported value and the value calculated with our fitting method. We note that, when using our fitting method, the result of van Dokkum et al. (2001) is very similar to the result presented in this paper. Splitting the van Dokkum et al. (2001) sample into low- and high-mass galaxies (at $M = 2 \times 10^{11} M_\odot$), we find $(-1.29 \pm 0.39) z$ for the high-mass galaxies and $(-1.94 \pm 0.34) z$ for the low-mass galaxies. We verified that this 2.5 $\sigma$ difference can be explained entirely by the fact that the galaxies are luminosity selected, not mass selected. Contrary to this study and the work of Treu et al. (2001), van Dokkum et al. (2001) did not select the galaxies by color, but by morphology and magnitude only.

Van Dokkum & Ellis (2003) find $(-1.25 \pm 0.25) z$. The difference in fitting method, mentioned above, applies to this study as well. Additionally, their transformation from $I$-band surface brightness to rest-frame $B$-band surface brightness, using $V - I$ colors, is uncertain. Namely, using $V - I$ is an extrapolation for galaxies at $z > 0.8$. This leads to a different value for the $M/L$ evolution if we use our method to fit to the individual galaxy data of the samples of van Dokkum et al. (2001) and van Dokkum & Ellis (2003): $(-1.68 \pm 0.13) z$, a value very similar to the result yielded by our sample. Using the biweight center to measure the evolution effectively gives very low weighting factors to outliers. If the two galaxies with the lowest $M/L$ in the high-z sample of van Dokkum et al. (2001) are omitted, we find $(-1.53 \pm 0.13) z$. The measurement accuracy, together with the uncertain transformation to rest-frame properties, can explain the remaining difference. The new sample described by van Dokkum & Ellis (2003) is too small to verify whether there is a trend with galaxy mass. There is one galaxy more massive than $M = 2 \times 10^{11} M_\odot$ with $\Delta \ln (M/L_B) = (-1.57 \pm 0.14) z$. As is the case for the sample of van Dokkum et al. (2001), the sample of van Dokkum & Ellis (2003) is not selected by color. This did not lead to faster evolution due to including blue early-type galaxies, which indicates that the color cuts in the other studies are adequately generous to avoid this bias.

G03 report a brightening in the rest-frame $B$ band of early-type galaxies by 2.4 mag at $z = 1$, which is derived from fitting a cubic spline to the offsets of the galaxies from the local FP. With our fitting method, we find $(-1.94 \pm 0.20) z$, using the data of the individual galaxies in that sample. This is consistent with the result presented in this paper. The difference between the value reported by G03 and our result using their data is caused by the three galaxies at redshifts $z > 0.9$. Our linear fit to their data and their fit agree well up to $z = 0.8$ (1.6 mag brightening). If we split the G03 sample into high- and low-mass galaxies, we find an evolution of $(-1.59 \pm 0.50) z$ for galaxies more massive than $M = 2 \times 10^{11} M_\odot$, and $(-2.05 \pm 0.17) z$ for less massive galaxies. For low-mass galaxies our result, $(-1.97 \pm 0.16) z$, and that of G03 agree well. The fast evolution for massive galaxies in their sample remains unexplained, but given the large uncertainty and the small number of objects, the inconsistency is only mild. We note, however, that contrary to the early-type galaxies in our sample, the early-type galaxies in the G03 sample do not show a correlation between $M/L$ and rest-frame $U - B$ color. One of the galaxies in the G03 sample serves as an example of the uncertainty. HST 14176+5226 has been observed spectroscopically before by both
Ohyama et al. (2002) and Treu & Koopmans (2004) and happens to be in the sample of lensing galaxies of Rusin et al. (2003) and van de Ven et al. (2003). The three independent spectroscopic velocity dispersions are \( \sigma = 222 \pm 8 \) (G03), 245 \( \pm 15 \) (Ohyama et al. 2002), and 224 \( \pm 15 \) km s\(^{-1}\) (Treu & Koopmans 2004), all corrected to the same aperture. Van de Ven et al. (2003) report a value of \( \sigma = 292 \pm 29 \) km s\(^{-1}\) as derived from the lensing model. This might be an indication that using dispersions derived from lensing models leads to slower evolution than using spectroscopic dispersions. Also, the consistency among the different spectroscopic dispersions shows that it remains unclear why G03 find somewhat faster evolution than the other authors. Repeated observation of the galaxies in the G03 sample may be illuminating.

We conclude that all the results found in the literature are mutually consistent, once differences in calculating and presenting the results have been taken into account.

6. CONCLUSIONS

We obtained ultradeep spectroscopy for 27 field early-type galaxies with redshifts 0.6 \( < z < 1.15 \). The offset of these high-redshift galaxies from the local FP is used as a measure of the evolution of \( M/L \) and as an age estimator.

The average evolution of the early-type galaxies in our sample is \( \Delta \ln (M/L_B) = (-1.75 \pm 0.16)z \). The value we find for galaxies in the primary sample, those galaxies satisfying all our selection criteria, is the same. The scatter in \( \Delta \ln (M/L_B) \) is large: 0.58. This shows that some galaxies must have high luminosity-weighted formation redshifts \( (z > 2) \), while others have formed a large fraction of their stars at redshifts 1 \( < z < 2 \). Emission lines in the spectra indicate that some galaxies show signs of ongoing star formation at the epoch of observation. This is in agreement with the presence of massive early-type galaxies with emission lines at \( z \sim 0.5 \) (Treu et al. 2002), although the galaxies with emission lines in our sample tend to have low masses.

We find a tight correlation between \( M/L \) and rest-frame color, which shows that the variation in \( M/L \) among the galaxies in the sample is intrinsic and due to differences in the stellar populations. The galaxies in our sample span a large range of masses. We find that low-mass galaxies have larger offsets from the local FP than high-mass galaxies. Because luminosity-selected samples are biased toward galaxies with low \( M/L \), this is a trend that is expected. We carefully analyze whether the observed correlation between mass and \( M/L \) can entirely be explained by this selection effect or not. We find that galaxies with masses \( M < 6 \times 10^{10} M_\odot \) are only included in our sample because they have low \( M/L \). For galaxies at \( z \sim 1 \) with masses larger than \( M = 6 \times 10^{10} M_\odot \), our sample is biased, but to a limited amount. Taking into account the selection effect, we exclude with high confidence that the distribution of mass and \( M/L \) of our \( z \sim 1 \) galaxy sample with masses \( M > 6 \times 10^{10} M_\odot \) is the same as that of the low-redshift field early-type galaxy sample taken from Bernardi et al. (2003), corrected for evolution. We do not claim that we have observed a change of the slope of the FP because we cannot exclude the possibility that our sample is drawn from a distribution that has the same slope as the Bernardi et al. (2003) sample, but with a scatter that is twice as large at \( z = 1 \). However, the outliers do not occur at the high-mass end \( (M \sim 10^{12} M_\odot) \) of our galaxy sample, but at lower masses, namely, \( M \sim 10^{11} M_\odot \). Therefore, our results show that the evolution of early-type galaxies is mass dependent, whether by an increase in the scatter at lower masses, or by systematic faster evolution of lower mass galaxies as compared to higher mass galaxies, or, which is the most natural explanation, by a combination of these effects. Assuming that the scatter has decreased from \( z = 1 \) to the present day by a factor of 2, we find that the bias-corrected \( M/L \) evolution of \( z \sim 1 \) early-type galaxies with masses \( M > 6 \times 10^{10} M_\odot \) is \( \Delta \ln (M/L_B) = (-1.43 \pm 0.16)z \).

Previous studies (Treu et al. 2001, 2002; van Dokkum et al. 2001; van Dokkum & Ellis 2003; G03) that claimed to have derived mutually exclusive results are in fact consistent with our results, if the same fitting method is applied to the different data sets. Particularly interesting is the consistency of our results with the results from a sample of lensing galaxies (Rusin et al. 2003; van de Ven et al. 2003), which also shows the mass dependence, even though this sample is not biased toward galaxies with low \( M/L \): because of the selection technique, the lens sample contains galaxies with typical \( M/L \) at a given mass. Our luminosity-limited sample is sensitive to outliers, which are present indeed. The combination of these independent results strengthens the evidence for mass-dependent evolution and the combined increase in both slope and scatter with redshift.

Bell et al. (2004) claim that the mass density of red sequence galaxies increases by at least a factor of 2 from \( z \sim 1 \) to the present. This is partly based on the observation that the luminosity density is constant out to \( z = 1 \). The \( M/L \) evolution of our galaxy sample implies an increase by a factor of 4 in the mass density. In the local universe, most of the mass density in early-type galaxies is accounted for by galaxies with a velocity dispersion of \( \approx 225 \) km s\(^{-1}\) (Kochanek et al. 2000) or a mass of \( \approx 3 \times 10^{11} M_\odot \). If those galaxies, which evolve somewhat slower, dominate the evolution of the mass density, the increase is slightly less \((3 - 3.5)\).

The correlation between mass and \( M/L \) has been observed in clusters as well (Wuyts et al. 2004), but this can entirely be explained by selection effects. Both our field sample and the cluster samples found in the literature are not strongly biased for galaxy masses \( M > 2 \times 10^{11} M_\odot \). When applying this mass cut, the evolution for cluster galaxies is \( \Delta \ln (M/L_B) = (-1.12 \pm 0.06)z \) and the evolution for field galaxies is \( \Delta \ln (M/L_B) = (-1.07 \pm 0.17)z \). Galaxies with masses comparable to the mass of \( L^* \) galaxies in the local universe \((\approx 3 \times 10^{11} M_\odot)\) have luminosity-weighted ages that imply formation redshifts \( z > 2 \), independent of environment. If progenitor bias is important, the luminosity-weighted age of typical early-type galaxies in the local universe can be considerably lower (van Dokkum & Franx 2001).

In hierarchical formation models, the predicted difference between the \( M/L \) of field and cluster galaxies is \( \Delta \ln (M/L_B) = 0.55 \), independent of redshift (van Dokkum et al. 2001). This large difference is related to the difficulty of constructing isolated galaxies without active star formation, i.e., to the lack of a mechanism that truncates star formation from within the galaxy. Our results rule out this prediction at the 99.6% confidence level, up to \( z = 1.1 \).

Our findings are consistent with downsizing (Cowie et al. 1996; Kodama et al. 2004). This idea is independently corroborated by other observations, such as the decrease of the mass of “E+A” galaxies with time (Tran et al. 2003), the lack of star formation in massive galaxies at redshifts \( z \leq 1 \) (De Lucia et al. 2004), the fossil record of star formation in local early-type galaxies (Thomas et al. 2005), and the claim of mass-dependent evolution of spiral galaxies (Ziegler et al. 2002; Böhm et al. 2004). The lack of age differences between field and cluster galaxies and the suggested mass-dependent evolution of
early-type galaxies show that individual properties of a galaxy, and not environment, play an important role in its formation. We have shown that rest-frame optical colors can be used to measure galaxy masses at high redshift. This can be regarded as a step toward accurately calibrating spectral energy distribution (SED) fitting as a mass estimator. Certainly including Spitzer photometry in the rest-frame near-infrared will provide tight correlations between dynamically derived $M/L$ and $M/L$ derived from SED fitting. Line strengths of absorption features in our high-S/N spectra will connect low-redshift fossil record studies from SED fitting. Using the mass calibration for high-redshift galaxies, the evolution of the mass density and the mass function can be determined from volume-limited samples. This will provide strong constraints on formation theories and the importance of progenitor bias.

We thank the referee for many useful comments, enhancing the quality and the readability of the work. We thank the ESO staff for their professional and effective assistance during the observations. We also thank the Lorentz center for its hospitality during various workshops and the Leidsch Kerkhoven-Bosscha Fonds for its financial support. P. v. D. acknowledges support from grant HST-AR-09541.02-A. G. I. acknowledges support for this research from NASA grant NAG5-7697.

Note added in proof.—T. Trev et al. (ApJ, 622, L5 [2005]) report results that are very similar to those presented here, although their interpretation of selection effects is different. It appears to us that the controversy described in § 5 has now ended.

REFERENCES

Alexander, D. M., et al. 2003, AJ, 126, 539
Bell, E., McIntosh, D. H., Katz, N., & Weinberg, M. D. 2003, ApJ, 149, 289
Bell, E. F., et al. 2004, ApJ, 608, 752
Bernardi, M., et al. 2003, AJ, 125, 1866
Bertin, E., & Arnouts, S. 1996, A&AS, 117, 393
Blakeslee, J. P., et al. 2003, ApJ, 596, L143
Böhm, A., et al. 2004, A&A, 420, 97
Brusa, A. G., & Charlot, S. 2003, MNRAS, 344, 1000
Cole, S., Lacey, S., Baugh, C. M., & Frenk, C. S. 2000, MNRAS, 319, 168
Coleman, G. D., Wu, C.-C., & Weedman, D. W. 1980, ApJS, 43, 393
Cowie, L. L., Songaila, A., Ha, F. M., & Cohen, J. G. 1996, AJ, 112, 839
De Lucia, G., et al. 2004, ApJ, 610, L77
Diaferio, A., Kauffmann, G., Balogh, M. L., White, S. D. M., Schade, D., & Ellingson, E. 2001, MNRAS, 323, 999
Djorgovski, S., & Davis, M. 1987, ApJ, 313, 59
Dressler, A., Lynden-Bell, D., Burstein, D., Davies, R. L., Faber, S. M., Terlevich, R., & Wegner, G. 1987, ApJ, 313, 42
Faber, S. M., Wegner, G., Burstein, D., Davies, R. L., Dressler, A., Lynden-Bell, D., & Terlevich, R. L. 1989, ApJ, 398, 763
Franx, M. 1993, PASP, 105, 1058
Gebhardt, K., et al. 2003, ApJ, 597, 239 (G03)
Giacconi, R., et al. 2002, ApJS, 139, 369
Giavalisco, M., et al. 2004, ApJ, 600, L93
Holden, B. P., et al. 2005, ApJ, 620, L83
Jørgensen, I., Franx, M., & Kjaergaard, P. 1995, MNRAS, 276, 1341
———. 1996, MNRAS, 280, 287
Kauffmann, G., & Charlot, S. 1998, MNRAS, 297, L23
Kelson, D. D., Illingworth, G. D., Franx, M., & van Dokkum, P. G. 2000, ApJ, 531, 184
Kochanek, C. S., et al. 2000, ApJ, 543, 131
Kodama, T., et al. 2004, MNRAS, 350, 1005
Krist, J. 1995, in ASP Conf. Ser. 77, Astronomical Data Analysis Software and Systems IV, ed. R. A. Shaw, H. E. Payne, & J. J. E. Hayes (San Francisco: ASP), 349

—T. Trev et al. (ApJ, 622, L5 [2005]) report results that are very similar to those presented here, although their interpretation of selection effects is different. It appears to us that the controversy described in § 5 has now ended.