Discovery of a Radio Relic in the Massive Merging Cluster SPT-CL J2023-5535 from the ASKAP-EMU Pilot Survey

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Abstract

The ASKAP-EMU pilot survey is a deep wide-field radio continuum survey designed to cover the entire southern sky and a significant fraction of the northern sky up to +30°. Here, we report a discovery of a radio relic in the merging cluster SPT-CL J2023-5535 at z = 0.23 from the ASKAP-EMU pilot 300 square degree survey (800–1088 MHz). The deep high-resolution data reveal a ∼2 Mpc scale radio halo elongated in the east–west direction, coincident with the intracluster gas. The radio relic is located at the western edge of this radio halo stretched ∼0.5 Mpc in the north–south orientation. The integrated spectral index of the radio relic within the narrow bandwidth is $\alpha_{800\text{MHz}}^{1088\text{MHz}} = -0.76 \pm 0.06$. Our weak-lensing analysis shows that the system is massive ($M_{200} = 1.04 \pm 0.36 \times 10^{15} M_\odot$) and composed of at least three subclusters. We suggest a scenario, wherein the radio features arise from the collision between the eastern and middle subclusters. Our discovery illustrates the effectiveness of the ASKAP-EMU survey in detecting diffuse emissions in galaxy clusters and when completed, the survey will greatly increase the number of merging cluster detections with diffuse radio emissions.

Unified Astronomy Thesaurus concepts: Galaxy clusters (584); Weak gravitational lensing (1797); Radio continuum emission (1340); Radio telescopes (1360); X-ray astronomy (1810)

1. Introduction

Large-scale diffuse radio emissions provide critical information for understanding galaxy cluster mergers. They can be classified into two broad categories: halos and relics. Radio halos are diffuse sources without distinct optical counterparts and are found in the central regions of merging clusters. Radio relics also do not have optical counterparts, but they are located in the cluster periphery and in general possess high levels of polarization. Although both radio halos and relics are indicators of cluster merger activities, radio relics have been considered a stronger constraint on the merger history because they can be used as direct probes of merger shocks (see reviews by Ferrari et al. 2008; Feretti et al. 2012; van Weeren et al. 2019).

The kinetic energy that is dissipated during cluster–cluster mergers can power the observed cluster-scale radio emission. However, the complex chain of physical mechanisms that leads to the acceleration of emitting particles and amplification of magnetic fields in the intracluster medium (ICM) are still poorly understood (e.g., Brunetti & Jones 2014). Giant radio halos are thought to originate from stochastic reacceleration induced by cluster merger turbulence (e.g., Brunetti et al. 2001; Petrosian 2001; Brunetti & Lazarian 2007; Miniati 2015); the contribution from secondary particles generated by the chain of hadronic collisions in the ICM has also been explored in the past (e.g., Dennison 1980; Blasi & Colafrancesco 1999) and more recently in combination with turbulent reacceleration models (e.g., Brunetti & Lazarian 2011; Pinzke et al. 2017). Radio relics are believed to originate from merger shocks (e.g., Enßlin et al. 1998). The original approach was based on the diffusive shock acceleration (DSA; Bell 1978; Drury 1983; Malkov & Drury 2001) of thermal electrons. However, the efficiency of acceleration at weak shocks in the ICM appears too low to reproduce the spectrum and luminosity of a large fraction of the observed radio relics (e.g., Botteon et al. 2020). One of the most popular modifications of this scenario is based on shock reacceleration of preexisting relativistic plasma (e.g., Kang & Ryu 2011; Pinzke et al. 2013; Kang & Ryu 2016), which has been supported in some cases by the connection between radio relics and active galactic nuclei (AGNs; e.g., Bonafede et al. 2014; van Weeren et al. 2017).

In order to better understand the origin of radio halos and relics, a current priority should be to increase the sample size. To date, there are ∼60 known radio relics. Because of the
cluster-to-cluster variation, the existing sample is too small to enable studies, where one can extract overarching principles. Among the upcoming concerted efforts, presently the Australian Square Kilometer Array Pathfinder\textsuperscript{17}-Evolutionary Map of the Universe\textsuperscript{18} (ASKAP-EMU; Norris et al. 2011) is the largest deep ($\sim 10 \, \mu Jy \text{ beam}^{-1}$), high-resolution ($\sim 10''$) radio continuum survey designed to cover the entire southern sky and a significant fraction of the northern sky up to $+30^\circ$. One of the scientific goals of the project is to enlarge the sample of clusters with diffuse radio emissions by at least two orders of magnitude.

In this study, we report the discovery of a radio relic in the massive merging cluster SPT-CL J2023-5535\textsuperscript{19} (hereafter CL2023 for brevity) at $z = 0.23$. The presence of the diffuse radio emission in CL2023 has been reported in C. Zheng et al. (2020, in preparation), who used the Murchison Widefield Array\textsuperscript{20} (MWA; Tingay et al. 2013), the Australia Telescope Compact Array\textsuperscript{21} (Frater et al. 1992), and the Molonglo Observatory Synthesis Telescope\textsuperscript{22} (MOST Mills 1981; Robertson 1991) data. However, the insufficient spatial resolution and the several bright neighboring radio point sources have prevented the earlier work from clearly resolving the halo and relic. In this paper, we also present our weak-lensing (WL) and X-ray analyses of CL2023 based on the archival Dark Energy Camera (DECam; Flaugher et al. 2015) and Chandra data, respectively, which enhance interpretation of the current discovery.

We adopt a $\Lambda$CDM cosmology with $H_0 = 70 \, \text{km s}^{-1} \text{Mpc}^{-1}$, $\Omega_m = 0.3$, and $\Omega_{\Lambda} = 0.7$. The angular size of $1'$ corresponds to a length scale of $\sim 223 \, \text{kpc}$ at the cluster redshift $z = 0.23$.

### 2. Observations

#### 2.1. Radio

ASKAP has 36 antennae, 34 of which are placed within a region of $2.3 \, \text{km diameter}$ while the outer four extend the baselines up to $\sim 6 \, \text{km}$. In all cases, as many as 36 antennae were used, and always the four outer antennae were included because their UV coverage was critical to the imaging quality. However, in a few cases, some of the inner antennae were omitted because of maintenance or hardware issues.

At the focus of each antenna is a phased array feed (PAF), which subtends a solid angle of $\sim 30$ square degrees. Each PAF consists of 192 dual-polarization receivers. A weighted sum of the outputs of groups of receivers form $30$ beams. Individual receivers, in general, contribute to more than one beam. Therefore, adjacent beams are not completely independent. The 30 beams together cover an area of $\sim 30$ square degrees on the sky.

The weights of the individual beams are initially calibrated by observing the Sun placed successively at the center of each beam, and then adjusting the weights for maximum signal-to-noise. However, a radiator at the vertex of each antenna or the On-dish Calibrator (ODC) enables the gain of each receiver to be calibrated. As a result, the solution initially obtained from

the Sun observation is modified using the ODC calibration and used to adjust the weights.

Before (or sometimes after) the observation of each target, the calibrator source 1934–638 is observed for 200 s at the center of each of the 30 beams, to provide bandpass and gain calibration. No further calibration, other than self-calibration, is performed during the observation.

The radio data used in this paper were taken in 2019 July from the ASKAP-EMU Pilot Survey, based on 10 h integration (an rms noise level of 25–35 $\mu Jy$ beam$^{-1}$) for Scheduling Block 9351, with a frequency range of 800–1088 MHz. The reduction was performed with the ASKAPsoft\textsuperscript{23} pipeline, using a multiscale CLEAN algorithm and two Taylor terms (T0 and T1), which allow production of maps at a fiducial frequency of 943 MHz (T0) and the corresponding spectral indices (T1/T0). A more extensive description of the Pilot Survey will be provided by R. Norris et al. (2020, in preparation).

In this paper we present images from both the original (Figure 1(A)) and diffuse-enhanced (Figure 1(B)) versions. The latter was created by first masking out bright ($>40$ mJy) compact sources from the original version, then smoothing the masked image with an FWHM = $25''$ Gaussian kernel, and finally combining the smoothed image (purple) with the original image (green). The resulting image (Figure 1(B)) makes it easy to visually separate diffuse emissions from compact sources.

#### 2.2. Optical

CL2023 was observed with the DECam mounted on the 4 m Blanco telescope at the Cerro Tololo Inter-American Observatory (PI: von der Linden). Table 1 summarizes the observations for the $g$, $r$, and $i$ filters that we retrieved from the NOAO archive\textsuperscript{24} for the current study. The Community Pipeline (Valdes et al. 2014) is used for the basic data reduction (i.e., overscan, bias, flat, etc.). The calibrated images were stacked into a single mosaic image for each filter using SCAMP\textsuperscript{25} and SWARP.\textsuperscript{26} We used the $i$-band image for our WL analysis because it provides the sharpest point-spread function (PSF). Intermediate PSF models were constructed for the individual exposures through principal component analysis (Jee et al. 2007) and stacked to obtain the final PSF model for shape measurement. Readers are referred to the descriptions in our previous papers for details (e.g., Jee & Tyson 2011; Jee et al. 2013; Finner et al. 2017). After applying our S/N, color, magnitude, and shape measurement error cuts, we obtain a source density of $\sim 11$ galaxies per square arcmin.

#### 2.3. X-Ray

CL2023 was observed with the Chandra X-ray observatory on 2014 March 30 (ObsId: 15108—PI: Jones, ACIS-I detector, VFaint Mode, 20.81 ks). The data were reduced using the CIAO 4.11 software with CALDB 4.8.3. We reprocessed the raw data using the chandra_repro script to produce a level 2 event file. We then created a broadband (0.5–7 keV) exposure-corrected image with the fluximage script.

\textsuperscript{17} https://www.atnf.csiro.au/projects/askap/index.html
\textsuperscript{18} https://www.emu-survey.org/
\textsuperscript{19} The cluster itself was discovered by the ROSAT-ESO Flux-Limited X-ray (REFLEX) galaxy cluster survey (B"ohringer et al. 2004).
\textsuperscript{20} http://www.mwatelescope.org/
\textsuperscript{21} https://www.narrabri.atnf.csiro.au/
\textsuperscript{22} https://astronomy.swin.edu.au/research/utmost/?page_id=32
\textsuperscript{23} https://www.atnf.csiro.au/computing/software/askapsoft/sdp/docs/current/index.html
\textsuperscript{24} http://archive1.dm.noao.edu/
\textsuperscript{25} https://www.astromatic.net/software/scamp
\textsuperscript{26} https://www.astromatic.net/software/swarp
For our X-ray temperature measurement, point sources were masked out using the \texttt{wavdetect} script and background flares were removed with the \texttt{degrade} script. We extracted grouped X-ray spectra with the \texttt{specextract} script in such a way that each bin has a minimum signal-to-noise ratio of 5. Then, we performed spectral fitting with the \texttt{XSPEC (v12.10.1f)} package and used the absorbed MEKAL plasma model (Kaastra & Mewe 1993; Liedahl et al. 1995) within the 1–5 keV energy band. The Galactic hydrogen density and the cluster metal abundance were assumed to be $N_H = 5.1 \times 10^{20}$ cm$^{-2}$ (Dickey & Lockman 1990) and 0.3 solar, respectively.

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**Figure 1.** Multivwavelength observations of CL2023. (A) ASKAP-EMU full-resolution ($12'' \times 11''$) image. The green pan-shape regions indicate the areas for the X-ray surface brightness analysis. The approximate location of the leading edge of the relic is marked with the red arc. (B) Composite radio image after both diffuse (purple) and compact sources (green) are enhanced (Section 2.1). The diffuse radio halo is elongated $\sim$1 Mpc in the east–west orientation and at the western edge of the halo is the radio relic. White contours denote the $3\sigma_n \times 2^n$ levels of the smoothed (FWHM = 25$''$) image, where $n = 0, 1, 2, 3, 4,$ and $\sigma_n = 10 \mu$ Jy beam$^{-1}$. The yellow arrow points at the "link" feature between the relic and nearby radio galaxy (Section 4). (C) Optical color-composite image with the overlays of ASKAP-EMU radio (cyan), Chandra X-ray (magenta), and galaxy isodensity contours (white). We use the DECam $g$, $r$, and $i$ filters to represent the intensities in blue, green, and red, respectively. A total of $\sim$1800 cluster member candidates are selected based on their 4000 Å break features. (D) Same as (C) except that we overlay the WL mass contours (white) over the optical color-composite image. The mass distribution composed of three substructures is consistent with the galaxy density distribution. The spectroscopically confirmed BCG coincides with the X-ray peak.

### Table 1

| Filter | Date     | $i_{\text{exp}}$ (s) | Seeing (arcseconds) | $m_{\text{lim}}$ |
|--------|----------|----------------------|---------------------|------------------|
| $g$    | 2016 July 7 | 4875                 | 1.46                | 26.0             |
| $r$    | 2016 July 7 | 3375                 | 1.33                | 25.8             |
| $i$    | 2016 July 6 | 3625                 | 0.78                | 25.0             |

**Note.**

* This is the 5$\sigma$ limiting magnitude for point sources.
3. Results

3.1. Detection of a Radio Relic and Halo

Cross-matching the galaxy clusters detected in the Planck SZ survey (Planck Collaboration et al. 2016) with the ASKAP-EMU radio continuum survey, we discovered a ∼0.5 Mpc relic associated with CL2023 shown in Figure 1, based on its spectral index, morphology, orientation, and location. Also, a clear ∼1 Mpc radio halo is detected eastward of the relic, which was originally reported by C. Zheng et al. (2020, in preparation) as a halo candidate, whose observations, however, could not separate the halo from the relic because of the insufficient resolution.

The halo follows the diffuse X-ray emission (Figure 1(C)) whose peak coincides with the BCG (Figure 1(D)). The similarity in morphology between radio halo and ICM is typical of radio halos (e.g., Feretti et al. 1997; Govoni et al. 2001a; van Weeren et al. 2019, for a review). In the context of turbulent reacceleration models the nonthermal components (particles and magnetic fields) are powered by the damping of the energy flux of large-scale turbulence that is generated by the dynamics and energy density of the ICM component leading to a morphological connection between the two components (e.g., see Brunetti & Jones 2014 for a review); the details of the spatial correlation are sensitive to the way turbulence is generated in the thermal background plasma and relativistic particles are accelerated and transported in that turbulence.

The relic, elongated in the north–south orientation, lies ∼0.5 Mpc west of the BCG. This location also corresponds to the western edge of the diffuse X-ray emission. Close inspection of the relic reveals that its southern edge is connected to a compact radio source via a faint “link” (yellow arrow in Figure 1(B)) reminiscent of the feature in A3411-12 (van Weeren et al. 2017). The compact radio source has an optical counterpart (Figure 2), which has a consistent color with that of the cluster red sequence. The location of this radio source also coincides with the western WL mass peak (see Section 3.2).

We determined the radio flux density of the relic to be $S_{\text{43MHz}} = 16.2 \pm 0.2$ mJy from the T0 image using a polygon aperture (see Figure 2) that traces the visual boundary of the feature. Dividing the T1 image by the T0 image gives an integrated spectral index of $\alpha_{\text{int}} = -0.76 \pm 0.06$. The extrapolated radio flux density of the relic at 1.4 GHz is $S_{\text{1.4GHz}} = 12.0 \pm 0.3$ mJy under the assumption that the spectral feature follows a single power law. The monochromatic luminosity and largest linear size of CL2023 are compared with those of other clusters retrieved from van Weeren et al. (2019) in Figure 3, which shows that CL2023 follows the relation.

To measure the radio flux density of the halo, we first masked out bright point sources (>40 mJy) and then replaced the fluxes with the in-halo average value. Within a polygon enclosing the halo, the flux density is $S_{\text{43MHz}} = 31.3 \pm 0.6$ mJy. The integrated spectral index of the halo and the extrapolated radio flux density at 1.4 GHz are $\alpha_{\text{int}} = -1.04 \pm 0.05$ and $S_{\text{1.4GHz}} = 20.8 \pm 0.3$ mJy, respectively, when we follow the same procedures used for the relic. Note that here we only quote statistical errors based on our rms noise measurement. The total errors should be larger when also systematic errors (e.g., flux scaling errors) are included. Full characterization of the systematic errors will be progressing over the next few years. We summarize the above radio property measurements in Table 2.

3.2. Cluster Galaxy and WL Mass Distribution

Detection of the radio relic suggests that CL2023 underwent a major merger. However, in order to reconstruct the merger scenario, we need to identify the cluster substructures contributing to the merger. To this end, we use both the galaxy and mass distributions.

Because no spectroscopic data of CL2023 are publicly available, we selected the cluster member candidates based on their 4000 Å break features. From the color–magnitude diagram, we chose a total of ∼1800 member candidates within the ∼13' × 13' region approximately centered at the BCG. We
3.3. Intrachannel Gas Properties

Using a circular (r = 146″ or 543 kpc) aperture centered at the X-ray peak, we determined the X-ray temperature of CL2023 to be $T_X = 7.97 \pm 0.79$ keV ($\Delta T_X = 0.91$). When the mass--temperature relation of Mantz et al. (2016) is employed, this temperature is converted to $M_{500} = 7.01 \pm 0.84 \times 10^{14} M_{\odot}$. Our X-ray mass estimate is consistent with the previous results. Tarrio et al. (2018) quote $M_{500} \sim 5 \times 10^{14} M_{\odot}$ based on their joint analysis of the Planck SZ and ROSAT X-ray data. The recent XMM-Newton study (Bulbul et al. 2019) reports $M_{500} = 6.49^{+0.81}_{-0.71} \times 10^{14} M_{\odot}$. These mass estimates based on X-ray data roughly agree with our WL-based result $M_{500} = 6.8 \pm 2.4 \times 10^{14} M_{\odot}$. However, given the clear indication of the on-going merger and invalidity of the single-halo assumption, we believe that the agreement is rather a coincidence.

Within the same r = 543 kpc aperture, the X-ray flux of CL2023 is $f_X = 21.2 \pm 0.5 \times 10^{-13}$ erg cm$^{-2}$s$^{-1}$, which is converted to a luminosity of $L_X = 3.41 \pm 0.10 \times 10^{44}$ erg s$^{-1}$. This Chandra luminosity is in good agreement with the ROSAT result $L_X = 3.26 \pm 0.88 \times 10^{44}$ erg s$^{-1}$ reported by Böhringer et al. (2004). The relation between the total radio luminosity of the halo $P_{1.4 \text{ GHz}} = 3.4 \pm 0.1 \times 10^{24}$ W Hz$^{-1}$ and the measured luminosity is consistent with the prediction from the $L_X - P_{1.4 \text{ GHz}}$ scaling relation of Feretti et al. (2012).

Although radio relics are believed to be tracers of merger shocks, only a few clusters have been shown to possess corresponding shock features in X-ray. Our Chandra data analysis suggests that CL2023 may belong to this rare class possessing a density jump across the relic. From the green “panda” regions depicted in Figure 1(A), we determined the density compression $C = 1.8 \pm 0.5$ as shown in Figure 4. From the same regions, we measured the temperatures of the pre- and post-shock regions to be $T_X = 7.3 \pm 3.3$ keV and $20 \pm 12$ keV, respectively. Given the current statistics, this temperature difference is insignificant.

4. Discussion

4.1. Too Flat Spectral Index for a Relic?

If particles are advected downstream and do not suffer from too strong adiabatic losses and reacceleration processes, the integrated spectral index is steeper than the injection spectral index by $\sim$0.5 (Brüggen et al. 2012; Brunetti & Jones 2014). Thus, the implied injection spectral index $\alpha_{\text{inj}} = -0.26 \pm 0.06$

$^{27}$ The total mass is greater than the sum of the three substructures because $r_{200}$ also increases.

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**Table 2**

| Halo    | Relic |
|---------|-------|
| $S_{\text{e40MHz}}$ (mJy) | 31.3 ± 0.6 | 16.2 ± 0.2 |
| $S_{\text{1.4GHz}}$ (mJy) | 20.8 ± 0.3 | 12.0 ± 0.3 |
| $P_{\text{1.4GHz}}$ ($10^{24}$ W Hz$^{-1}$) | 3.4 ± 0.1 | 1.8 ± 0.1 |
| Spectral Index ($\alpha$) | $-1.04 \pm 0.05$ | $-0.76 \pm 0.06$ |

**Note.**

$^a$ Flux densities at 1.4 GHz are extrapolated assuming a power law.

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**Figure 4.** X-ray surface brightness profile across the relic. Black crosses are the data points (see Figure 1(A) for the selected regions) while the best-fit broken power-law model based on PROFIT v1.5 (Eckert et al. 2011) is shown in blue. The red dashed line indicates the location of the relic boundary shown in Figure 1(A).

**Table 3**

| Substructure | $M_{500}^a$ ($10^{14} M_{\odot}$) | Peak Significance$^b$ ($\sigma$) |
|-------------|---------------------------------|-------------------------------|
| East        | 2.6 ± 1.6                       | 3.6                           |
| Center      | 3.5 ± 1.7                       | 5.0                           |
| West        | 1.5 ± 1.2                       | 3.0                           |

**Notes.**

$^a$ We obtain masses by simultaneously fitting three NFW profiles.

$^b$ Mass peak significances are measured from the $2'$ aperture by dividing the convergence by the rms map derived from our bootstrapping analysis.
of the CL2023 relic (Section 3.1) is significantly flatter than the theoretical upper limit \( \alpha_{\text{inj}} < -0.5 \) allowed in the DSA model. According to the recent review by van Weeren et al. (2019), integrated spectral indices of the relics are within the range \(-1.5 < \alpha_{\text{int}} < -1.0\). Therefore, taken at face value, the \( \alpha_{\text{int}} \) value of CL2023 is unusual. However, one must remember that we derived \( \alpha_{\text{int}} \) from the narrow bandwidth (800–1088 MHz), which has yet to be verified by observations at other frequencies. Note that the spectral index of the radio relic could be biased toward a steeper value because the feature blends into the halo, which has a relatively steep spectral index.

To avoid any significant contamination from the radio halo, we measured the flux density of the relic from a high-resolution image (see Figure 2).

In order to examine a consistency, we retrieved the archival MOST data (843 MHz), which has a beam size of 45" × 54" at the location of the relic. We degraded the resolution of our ASKAP-EMU image to match this resolution and derive an aperture correction factor of 1.37 for the same polygon aperture. With this aperture correction, we obtained a flux density of 17.9 ± 0.3 mJy, which gives a spectral index of \(-0.89 ± 0.19\), consistent with the ASKAP-EMU measurement \(-0.76 ± 0.06\).

Similarly to CL2023, flat integrated spectral indices are observed in some other clusters. For example, the radio relics of A2256 (van Weeren et al. 2012; Trasatti et al. 2015) and A3667 (Hindson et al. 2014) have spectral indices of \( \alpha_{\text{int}} = -0.83 ± 0.03 \) and \(-0.9 ± 0.1\), respectively.

If the flatness is confirmed, it would imply that the cooling time of the downstream electrons is significantly longer than the age of the shock inferred from the merger scenario (Section 4.2). This may happen if the relic originates from the reacceleration of a cloud of seed relativistic plasma (e.g., Kang et al. 2012). In this case the observed thickness of the relic traces the scale of the preexisting plasma and the role of the plasma aging would be insignificant, provided that the shock has just crossed the cloud and that the shock crossing time is much shorter than the electron cooling time.

As mentioned in Section 3.1, the radio image hints at a possible presence of a "bridge" connecting the relic to the nearby radio galaxy (Figure 2). If a future study reveals a spectral index gradient along the bridge, CL2023 might become one of the strong cases supporting the reacceleration scenario as previously shown by A3411-12 (van Weeren et al. 2017; Andrade-Santos et al. 2019).

### 4.2. Merger Scenario

Both location and orientation of radio relics provide constraints on merger scenarios. The current CL2023 relic orientation suggests that the merger might be happening in the east–west direction, which implies that the relic is the result of the collision between the middle and eastern subclusters. Under the assumption that the shock was generated at the impact and has been propagating to the west with nearly the same speed as the collision speed in the plane of the sky, its location (detected by the relic) can be used as an indicator of the time-since-collision (TSC). Here, we present our estimates of the TSC in two ways. In one method, we infer the collision velocity using the timing argument (Sarazin 2002). In the other, we use our Mach number measurements.

The timing argument (based on the assumption that the two clusters freefall to each other from an infinite separation) gives a relative velocity of \( \sim 2000 \text{ km s}^{-1} \) at the separation \( d \sim 1 \text{ Mpc} \). If we further assume that this velocity is representative of the impact velocity, the current \( \sim 0.5 \text{ Mpc} \) separation of the relic from the central substructure implies a TSC of \( \sim 0.3 \text{ Gyr} \).

The density compression \( C = 1.8 ± 0.5 \) (Section 3.3) corresponds to the Mach number \( M = 1.6 ± 0.5 \) under the Rankine–Hugoniot shock conditions. In order to derive a collision speed, we need to compute the sound speed \( c_s \), which is estimated to be \( c_s \sim 1300 \text{ km s}^{-1} \) from the preshock temperature \( \sim 7 \text{ keV} \). The resulting collision speed is \( \sim 2000 \text{ km s}^{-1} \), in good agreement with the value from the timing argument.

As mentioned in Section 4.1, we cannot obtain an injection spectral index \( \alpha_{\text{inj}} \) from the integrated spectral index \( \alpha_{\text{int}} \) under the stationary shock conditions. Assuming that \( \alpha_{\text{int}} \) is a lower limit on \( \alpha_{\text{inj}} \), we can convert \( \alpha_{\text{inj}} \geq -0.76 \) to \( M \geq 2.9 \), which in turn corresponds to TSC \( \leq 0.2 \text{ Gyr} \).

In summary, although we believe that the above TSC estimates should be refined with future studies, we note that the current data set of CL2023 indicates that the merger shock might have originated from a recent merger (0.2–0.3 Gyr).

### 5. Conclusions

From the deep high-resolution ASKAP-EMU pilot 300 square degree survey, we discovered a \( \sim 0.5 \text{ Mpc} \) scale radio relic in the massive galaxy cluster SPT-CL J2023–5535 at \( z = 0.23 \). We also confirmed the existence of the \( \sim 1 \text{ Mpc} \times 0.5 \text{ Mpc} \) radio halo previously reported as a halo candidate.

Our study with the multiwavelength data including Chandra and DECam shows that (1) the radio halo coincides with the intracluster gas, (2) the cluster is composed of three subclusters, and (3) across the relic there is a hint of density jump in X-ray. Based on these results, we suggest that the cluster is a post-merger system, where the middle and eastern subclusters might have suffered a major collision 0.2–0.3 Gyr ago. The cluster may belong to a rare class of radio relics clusters, where the integrated spectral indices of the relics are flatter than the test-particle DSA limit. It possibly indicates the presence of preaccelerated fossil electrons from the neighboring radio galaxy that are reaccelerated because of the merger shock.

There are well-known theories regarding particle acceleration by merger-induced turbulence for radio halos and shock (re)acceleration for radio relics. However, the physics of these mechanisms in the ICM is still poorly known. Multiwavelength observations in this framework provide a fundamental guide for the theory. With only \( \sim 60 \) known radio relic systems to date, clearly one outstanding difficulty is the small sample size. Fortunately, a few giant radio surveys with state-of-the-art telescopes (e.g., SKA, LOFAR, etc.) are planned for the coming decade. The ASKAP-EMU survey, as one important program, will greatly increase the sample size by at least two orders of magnitude. The current study based on its pilot 300 square degree data demonstrates its tremendous potential when the full survey becomes available and supported by other multiwavelength data.

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**References**

Andrade-Santos, F., van Weeren, R. J., di Gennaro, G., et al. 2019, ApJ, 887, 31  
Bell, A. R. 1978, MNRAS, 182, 147  
Blasi, P., & Colafrancesco, S. 1999, ApPh, 12, 169  
Böhringer, H., Schuecker, P., Guzzo, L., et al. 2004, A&A, 425, 367  
Bonafede, A., Intema, H. T., Brüggen, M., et al. 2014, ApJ, 785, 1  
Boitone, A., Brunetti, G., Ryu, D., & Roh, S. 2020, A&A, 634, A64  
Brüggen, M., Bykov, A., Ryu, D., et al. 2012, SSRv, 166, 187  
Brunetti, G., & Jones, T. W. 2014, JMPD, 23, 1430007-98  
Brunetti, G., & Lazarian, A. 2007, MNRAS, 378, 245  
Brunetti, G., & Lazarian, A. 2011, MNRAS, 410, 127  
Brunetti, G., Setti, G., Feretti, L., & Giovannini, G. 2001, MNRAS, 320, 365  
Bulbul, E., Chiu, I.-N., Mohr, J. J., et al. 2019, ApJ, 871, 50  
Dennison, B. 1980, ApJL, 239, L93  
Dickey, J. M., & Lockman, F. J. 1990, ARA&A, 28, 215  
Drury, L. O. C. 1983, RPPh, 46, 973  
Eckert, D., Molendi, S., & Paltani, S. 2011, A&A, 526, A79  
Enßlin, T. A., Birnbaum, P. L., Klein, U., & Kohle, S. 1998, A&A, 332, 395  
Feretti, L., Boehringer, H., Giovannini, G., & Neumann, D. 1997, A&A, 317, 432  
Feretti, L., Giovannini, G., Govoni, F., & Murgia, M. 2012, A&ARv, 20, 54  
Ferrari, C., Govoni, F., Schindler, S., Bykov, A. M., & Rephaeli, Y. 2008, SSRv, 134, 93  
Finner, K., Jee, M. J., Golovich, N., et al. 2017, ApJ, 851, 46  
Fischer, P., & Tyson, J. A. 1997, AJ, 114, 14  
Flaugher, B., Diehl, H. T., Honscheid, K., et al. 2015, AJ, 150, 150  
Frater, R. H., Brooks, J. W., & Whiteoak, J. B. 1992, JEEEA, 12, 103  
Govoni, F., Enßlin, T. A., Feretti, L., & Giovannini, G. 2001a, A&A, 369, 441  
Hindson, L., Johnston-Hollitt, M., Hurley-Walker, N., et al. 2014, MNRAS, 445, 330  
Jee, M. J., Blakeslee, J. P., Sirianni, M., et al. 2007, PASP, 119, 1403  
Jee, M. J., & Tyson, J. A. 2011, PASP, 123, 596  
Jee, M. J., Tyson, J. A., Schneider, M. D., et al. 2013, ApJ, 765, 74  
Kaastra, J. S., & Mewe, R. 1993, A&AS, 97, 443  
Kaiser, N., & Squires, G. 1993, ApJ, 404, 441  
Kang, H., & Ryu, D. 2011, ApJ, 734, 18  
Kang, H., & Ryu, D., 2016, ApJ, 823, 13  
Kang, H., Ryu, D., & Jones, T. W. 2012, ApJ, 756, 97  
Liedahl, D. A., Osterheld, A. L., & Goldstein, W. H. 1995, ApJL, 438, L115  
Malkov, M. A., & Drury, L. O. C. 2001, RPPh, 64, 429  
Mantz, A. B., Allen, S. W., Morris, R. G., et al. 2016, MNRAS, 463, 3582  
Mills, B. Y. 1981, PASA, 4, 156  
Miniati, F. 2015, ApJ, 800, 60  
Navarro, J. F., Frenk, C. S., & White, S. D. M. 1997, ApJ, 490, 493  
Norris, R. P., Hopkins, A. M., Afonso, J., et al. 2011, PASA, 28, 215  
Petrosian, V. 2001, ApJ, 557, 560  
Pinzke, A., Oh, S. P., & Pfrommer, C. 2013, MNRAS, 435, 1061  
Pinzke, A., Oh, S. P., & Pfrommer, C. 2017, MNRAS, 465, 4800  
Planck Collaboration, Ade, P. A. R., Aghanim, N., et al. 2016, A&A, 594, 27  
Robertson, J. G. 1991, A&Ph, 44, 729  
Sarazin, C. L. 2002, ASSI, 272, 1  
Tarrío, P., Melin, J. B., & Arnaud, M. 2018, A&A, 614, A82  
Tingay, S., Goeke, R., Bowman, J., et al. 2013, PASA, 30, 7  
Trasatti, M., Akamatsu, H., Lovisari, L., et al. 2015, A&A, 575, A45  
Valdes, F., Gruendl, R. & DES Project 2014, in ASP Conf. Series 485, Astronomical Data Analysis Software and Systems XXIII, ed. N. Manset & P. Forshay (San Francisco, CA: ASP), 379  
van Weeren, R., Andrade-Santos, F., Dawson, W., et al. 2017, NatAs, 1, 0005  
van Weeren, R. J., de Gasperin, F., Akamatsu, H., et al. 2019, SSRv, 215, 16  
van Weeren, R. J., Röttgering, H. J. A., Rafferty, D. A., et al. 2012, A&A, 543, A43