Behavior of Einstein-Rosen Waves at null infinity

Abhay Ashtekar\textsuperscript{1}, Jiří Bičák\textsuperscript{2}, and Bernd G. Schmidt\textsuperscript{3}

\textsuperscript{1} Center for Gravitational Physics and Geometry
Department of Physics, Penn State, University Park, PA 16802, USA
\textsuperscript{2} Department of Theoretical Physics, Charles University
V Holešovičkách 2, 180 00 Prague 8, Czech Republic
\textsuperscript{3} Max-Planck-Institut für Gravitationsphysik,
Schlaatzweg 1, 14473 Potsdam, Germany

Abstract

The asymptotic behavior of Einstein-Rosen waves at null infinity in 4 dimensions is investigated in all directions by exploiting the relation between the 4-dimensional space-time and the 3-dimensional symmetry reduction thereof. Somewhat surprisingly, the behavior in a generic direction is better than that in directions orthogonal to the symmetry axis. The geometric origin of this difference can be understood most clearly from the 3-dimensional perspective.

I. INTRODUCTION

Although the literature of Einstein-Rosen waves is quite rich (see, e.g., the references listed in the companion paper \cite{1}) it appears that there is only one article that discusses the asymptotic behavior of these waves at infinity in 4-dimensions: the paper by Stachel \cite{2} written already in the sixties. Moreover, even in this work, Stachel deals solely with the directions orthogonal to the axis of symmetry, i.e., to the $\partial/\partial z$ Killing field. The purpose of this note is to analyze the asymptotic structure in all directions.

Since these space-times admit a translational Killing field, one would expect them not to be asymptotically flat. This is precisely what Stachel discovered in directions orthogonal to the symmetry axis. Somewhat surprisingly, however, we will find that the fall-off is much better in generic directions. Indeed, if one restricts oneself to the “time-symmetric” situation, one finds that, in all other directions, curvature peels normally and a regular null infinity, $I$, exists. In fact, all radiation is concentrated along the two generators in which the null geodesics orthogonal to the symmetry axis meet $I$. In other directions, there is curvature but no radiation. If one goes beyond the “time-symmetric” case, the behavior is not as nice; $I$ may have a logarithmic character \cite{3,4}. That is, the metric does admit Bondi-type expansions but in terms of $r^{-3}\ln r$. Nonetheless, even this behavior is better than the one encountered in the directions orthogonal to the symmetry axis.
The key idea behind our analysis is to exploit the relation between 4-dimensional Einstein-Rosen waves and the associated 3-dimensional geometry on the manifold of orbits of the translational Killing field. Since the translational Killing field has been “factored out” in the passage to 3 dimensions, the 3-dimensional space-time is asymptotically flat at null infinity\[1\] and, as we will see, also admits a regular time-like infinity. To analyze the behavior of the 4-dimensional metric, we can draw on this 3-dimensional information. We will find that the behavior at \(I\) in generic directions in 4 dimensions is dictated by the behavior of various fields at time-like infinity in 3 dimensions. Since this 3-dimensional time-like infinity is regular, the behavior in generic directions in 4 dimensions is better than what one might naively expect.

The plan of the paper is as follows. In Sec. II, we shall present the 3-dimensional structure. In Sec. III, we will use this structure to investigate 4-dimensional null infinity. Appendix A spells out the relation between 3- and 4-dimensional curvatures.

**II. 3-DIMENSIONAL DESCRIPTION**

This section is divided into three parts. In the first, we briefly recall the symmetry reduction procedure and apply it to obtain the 3-dimensional equations governing Einstein-Rosen waves. (For details, see \[1\].) This procedure reduces the task of finding a 4-dimensional Einstein-Rosen wave to that of finding a solution to the wave equation on 3-dimensional *Minkowski* space. In the second part, we analyze the asymptotic behavior of these solutions to the wave equation at time-like infinity of the 3-dimensional Minkowski space. In the third part, we combine the results of the first two to analyze the asymptotic behavior of the 3-dimensional metric associated with Einstein-Rosen waves at its time-like infinity. We show that this time-like infinity is regular. Although this result is not needed directly for our main result, it is included because it complements the general analysis of 3-dimensional null infinity presented in \[1\].

A. Symmetry reduction

Recall first that the metric of a vacuum space-time with two commuting, hypersurface orthogonal space-like Killing vectors can always be written locally as \[5\]

\[
ds^2 = e^{2\psi}dz^2 + e^{2(\gamma-\psi)}(-dt^2 + d\rho^2) + \rho^2 e^{-2\psi}d\phi^2 ,
\]

where \(\rho\) and \(t\) (the “Weyl canonical coordinates”) are defined invariantly and \(\psi = \psi(t,\rho), \gamma = \gamma(t,\rho)\). (Here, some of the field equations have been used.) Einstein-Rosen waves have cylindrical symmetry; the Killing field \(\partial/\partial z\) is translational and \(\partial/\partial \phi\) is rotational and keeps a time-like axis fixed. Then the coordinates used in (2.1) are unique up to a translation \(t \mapsto t + a\).
The 3-manifold is obtained by quotienting the 4-dimensional space-time by the orbits of the \( \partial/\partial z \) Killing field and is thus co-ordinatized by \( t, \rho \) and \( \phi \). The 4-metric naturally induces a 3-metric \( d\bar{\sigma}^2 \) on this manifold and the 4-dimensional Einstein’s equations can be expressed on the 3-manifold as a system of coupled equations involving the induced 3-metric and the norm of the Killing field \( \partial/\partial z \), which, from the 3-dimensional perspective can be regarded as a (scalar) matter field. It is well-known (see \[1,6,7\]), however, that the field equations simplify considerably if we rescale the induced 3-metric \( d\bar{\sigma}^2 \) by \( e^{2\psi} \), the square of the norm of the Killing field, i.e., in terms of the 3-metric
\[
d\sigma^2 = g_{ab}dx^a dx^b = e^{2\gamma}(-dt^2 + d\rho^2) + \rho^2 d\phi^2 .
\] (2.2)
The 4-dimensional vacuum equations are then equivalent to the set (cf. Eqs. (2.12)–(2.15) in the preceding paper \[4\]):
\[
\gamma'' - \ddot{\gamma} + \rho^{-1}\gamma' = 2\dot{\psi}^2 ,
\] (2.3)
\[
-\gamma'' + \ddot{\gamma} + \rho^{-1}\gamma' = 2\psi'^2 ,
\] (2.4)
\[
\rho^{-1}\dot{\gamma} = 2\dot{\psi}\psi' ,
\] (2.5)
and
\[
-\ddot{\psi} + \psi'' + \rho^{-1}\psi' = 0 ,
\] (2.6)
on the 3-manifold, where the dot and the prime denote derivatives with respect to \( t \) and \( \rho \) respectively. The last equation is the wave equation for the non-flat 3-metric (2.2) as well as for the flat metric obtained by setting \( \gamma = 0 \). This is a key simplification for it implies that the equation satisfied by the matter source \( \psi \) decouples from the equations (2.3)–(2.5) satisfied by the metric. Furthermore, these latter equations reduce simply to:
\[
\gamma' = \rho(\dot{\psi}^2 + \psi'^2) ,
\] (2.7)
\[
\dot{\gamma} = 2\rho\dot{\psi}\psi' .
\] (2.8)
Thus, we can first solve for the axi-symmetric wave equation (2.6) for \( \psi \) on Minkowski space and then solve (2.7) and (2.8) for \( \gamma \) – the only unknown metric coefficient – by quadratures. (Note that (2.7) and (2.8) are compatible because their integrability condition is precisely (2.6).)

B. Asymptotic behavior of scalar waves

In this subsection we will focus on the axi-symmetric wave equation in 3-dimensional Minkowski space and analyze the behavior of its solutions \( \psi \) near timelike infinity of Minkowski space. (For behavior at null infinity, see \[3\].)
We begin with an observation. The “method of descent” from the Kirchhoff formula in 4 dimensions gives the following representation of the solution of the wave equation in 3 dimensions, in terms of Cauchy data \( \Psi_0 = \psi(t = 0, x, y), \Psi_1 = \psi_s(t = 0, x, y) \):

\[
\psi(t, x, y) = \frac{1}{2\pi} \frac{\partial}{\partial t} \int_{S(t)} \frac{\Psi_0(x', y') dx' dy'}{[t^2 - (x - x')^2 - (y - y')^2]^{1/2}} + \frac{1}{2\pi} \int_{S(t)} \frac{\Psi_1(x', y') dx' dy'}{[t^2 - (x - x')^2 - (y - y')^2]^{1/2}},
\]

(2.9)

where \( S \) is the disk

\[(x - x')^2 + (y - y')^2 \leq t^2\]

in the initial Cauchy surface (see e.g. [8]). We will assume that the Cauchy data are axially symmetric and of compact support.

In the preceding paper [1] (see Eq. (2.23)) we have shown that on each null hypersurface \( u = t - \rho = \text{const} \) the solution (2.9) can be expanded in the form

\[
\psi(u, \rho) = \frac{1}{\sqrt{\rho}} \left( f_0(u) + \sum_{k=1}^{\infty} \frac{f_k(u)}{\rho^k} \right),
\]

(2.10)

where the coefficients in this expansion are determined by integrals over the Cauchy data. This is the behavior of \( \psi \) at null infinity \( I^+ \).

Let us now investigate the behavior of the solution (2.9) near time-like infinity \( i^+ \) of the 3-dimensional Minkowski space. Setting

\[
t = U + \kappa \rho, \quad \kappa > 1,
\]

(2.11)

we wish to find \( \psi \) for \( \rho \to \infty \) with \( U \) and \( \kappa \) fixed. For large enough \( \rho \) the region of integration is contained in the cone. Hence we have to perform the derivative in (2.9) only in the integrand. We obtain

\[
2\pi \psi(t, \rho) =
\frac{\kappa \rho + U}{[(\kappa^2 - 1)\rho^2]^{3/2}} \int_0^\infty \int_0^{2\pi} \Psi_0 \rho' d\rho' d\phi' \left[ 1 + \frac{2(\kappa U + \rho' \cos \phi')}{\kappa^2 - 1} \frac{1}{\rho} + \frac{U^2 - \rho^2}{\kappa^2 - 1} \frac{1}{\rho^2} \right]^{-3/2}
+ \frac{1}{[(\kappa^2 - 1)\rho^2]^{1/2}} \int_0^\infty \int_0^{2\pi} \Psi_1 \rho' d\rho' d\phi' \left[ 1 + \frac{2(\kappa U + \rho' \cos \phi')}{\kappa^2 - 1} \frac{1}{\rho} + \frac{U^2 - \rho^2}{\kappa^2 - 1} \frac{1}{\rho^2} \right]^{-1/2}.
\]

(2.12)

The integrand can again be expanded in \( \rho^{-1} \) (or \( t^{-1} \)) but the leading term is \( \rho^{-1} \). By contrast, at null infinity of the 3-dimensional space-time, \( \psi \) falls-off only as \( \rho^{-1/2} \) (see Eq. (2.10) and [1] for details). We will see that it is this difference that makes the behavior of the 4-metric along generic directions better than that along directions orthogonal to the symmetry axis.
The explicit expressions of the first few terms in the expansion of $\psi$ is given by:

$$\psi = \frac{L}{(\kappa^2 - 1)^{3/2}} \left[ -\frac{\kappa}{\rho^2} + \frac{(2\kappa^2 + 1)U}{\kappa^2 - 1} \frac{1}{\rho^3} + O(\rho^{-4}) \right]$$

$$+ \frac{J}{(\kappa^2 - 1)^{1/2}} \left[ \frac{1}{\rho} - \frac{\kappa U}{\kappa^2 - 1} \frac{1}{\rho^2} + O(\rho^{-3}) \right], \quad (2.13)$$

where

$$L = \int_0^\infty \Psi_0(\rho') \rho' d\rho', \quad J = \int_0^\infty \Psi_1(\rho') \rho' d\rho'. \quad (2.14)$$

By expressing $\rho$ in terms of $t$ using (2.11), we may rewrite (2.13) as a series in $t^{-1}$:

$$\psi = \frac{L}{(\kappa^2 - 1)^{3/2}} \left[ -\frac{\kappa^3}{t^2} + \frac{3\kappa^3 U}{\kappa^2 - 1} \frac{1}{t^3} + O(t^{-4}) \right]$$

$$+ \frac{J}{(\kappa^2 - 1)^{1/2}} \left[ \frac{\kappa}{t} - \frac{\kappa U}{\kappa^2 - 1} \frac{1}{t^2} + O(t^{-3}) \right]. \quad (2.15)$$

The last formula is meaningful also for $\rho = 0$ in the limit $\kappa \to \infty$:

$$\psi = -\frac{L}{t^2} + \frac{J}{t} + O(t^{-3}). \quad (2.16)$$

The same result can be obtained from (2.9) directly. This concludes our discussion of the asymptotic behavior of $\psi$ near time-like infinity $i^\pm$.

We will conclude this sub-section with three remarks.

First, the explicit representation (2.9) of the solution in terms of Cauchy data allows us to make the interesting observation that the solution is actually analytic in its space-time dependence for all points for which the data are within the past null cone. To show that all solutions with data of compact support are also analytic in a neighborhood of future time-like infinity $i^+$ we have to use conformal rescaling techniques. Let

$$d\sigma^2 = -dt^2 + dx^2 + dy^2$$

be the metric of 3-dimensional Minkowski space. The conformal factor

$$\Omega = (t^2 - x^2 - y^2)^{-1}$$

defines, by the rescaling $d\tilde{\sigma}^2 = \Omega^2 d\sigma^2$, again a flat space-time

$$d\tilde{\sigma}^2 = \Omega^2 d\sigma^2 = -d\tilde{t}^2 + d\tilde{x}^2 + d\tilde{y}^2,$$

where the coordinates $\tilde{t}, \tilde{x}, \tilde{y}$ are defined by the relations (inversion)

$$\tilde{t} = \frac{t}{t^2 - x^2 - y^2}, \quad \tilde{x} = \frac{x}{t^2 - x^2 - y^2}, \quad \tilde{y} = \frac{y}{t^2 - x^2 - y^2}. \quad (2.20)$$
The 3-dimensional scalar wave equation has the following behavior under this conformal re-scaling:

\[ \nabla^2 \psi = 0 \implies \tilde{\nabla}^2 \tilde{\psi} = 0, \quad \tilde{\psi} = \Omega^{-1/2} \psi. \quad (2.21) \]

From the above consideration we know that a solution \( \psi \) with data of compact support is analytic for points within and on the future light cone of the point \( \tilde{t} = a, \tilde{x} = \tilde{y} = 0 \), where the value of \( a \) is dictated by the support of the data. Moreover the series (2.10) is also analytic in \( \nu = u + 2\rho \) because of the converging expansion in \( \rho^{-1} \). Hence after the inversion we have a solution \( \tilde{\psi} \) which is analytic on the extended null cone. Therefore it is analytic in a domain which includes a neighborhood of \( i^+ \).

The second remark concerns the asymptotic behavior of \( \psi \), regarded as a solution to the wave equation in 4 dimensions. More precisely, let us set

\[ F(t, x, y, z) = \psi(t, x, y); \quad (2.22) \]

\( F \) is independent of \( z \). How does this solution behave at null infinity of 4-dimensional Minkowski space? The null geodesics in a hypersurface \( z = \text{const} \) are also null geodesics in 4-space and \( F = \psi \) along these curves. Now, a solution of the 4-dimensional wave equation is well behaved at null infinity if it falls off as \( r^{-1} \) (where \( r \) is the standard radial coordinate). Since the field \( \psi \) falls off only as \( \rho^{-1/2} \) at null infinity in 3-dimensions, the solution \( F \) fails to define a finite radiation field at null infinity in these directions. For null lines not contained in \( z = \text{const} \) surfaces, on the other hand, the situation is entirely different. Because such null lines project onto time-like lines in \( z = \text{const} \), the fall-off behavior is much better and from (2.15) we obtain the \( r^{-1} \) decay, necessary for the radiation field to exist. Thus, in terms of a 4-dimensional conformal rescaling, the rescaled field of \( F \) will be well-defined on 4-dimensional null infinity except for the two null generators determined by the \( \partial/\partial z \)-Killing vector. We will see in Section III that this behavior is the key to the understanding of the asymptotics of 4-dimensional axi-symmetric space-times with a further \( \partial/\partial z \)-Killing vector.

Finally, we wish to point out that the main results obtained in this section continue to hold also for general data of compact support which are not necessarily axi-symmetric. In particular, asymptotic forms like (2.13) and (2.15) hold where, however, the coefficients depend on \( \phi \). The assumption of compact support can also be weakened to allow data which decay near spatial infinity sufficiently rapidly so that we still obtain solutions smooth at null and time-like infinities. This is the case, for example, with the Weber–Wheeler–Bonnor pulse discussed in the following section.

C. Asymptotic behavior of the metric

We now combine the results of the previous two sub-sections. Recall from Eq. (2.2) that the 3-dimensional metric \( d\sigma^2 \) has a single unknown coefficient, \( \gamma(t, \rho) \), which is determined by the solution \( \psi(t, \rho) \) to the wave equation in Minkowski space
(obtained simply by setting \( \gamma = 0 \)). The asymptotic behavior of \( \psi(t, \rho) \) therefore determines that of the 3-metric.

At null infinity \( I \), the asymptotic behavior (2.10) of \( \psi \) implies that \( \gamma \) has the form (see Eq.(2.32) in [1]):

\[
\gamma = \gamma_0 + \int_{-\infty}^{u} \left( -2(\dot{f}_0(u))^2 + \sum_{k=1}^{\infty} \frac{g_k(u)}{\rho^k} \right) du .
\]  

(2.23)

We now wish to determine the metric at \( i^+ \). In the last sub-section we found the asymptotic form of \( \psi \) at \( i^+ \), more specifically, at \( \rho \to \infty \) (or \( t \to \infty \) with \( U = t - \kappa \rho \), \( \kappa > 1 \), fixed (see (2.11) –(2.16)). In order to get the asymptotic forms of \( \gamma \), we first express the field equations (2.7) and (2.8) for \( \gamma \) in terms of \( U \) and \( \rho \):

\[
\gamma_{,U} = 2\rho \psi_{,U}(\psi_{,\rho} - \kappa \psi_{,U}) \, ,
\]  

(2.24)

\[
\gamma_{,\rho} = \rho \left[ \psi_{,\rho}^2 + (1 - \kappa^2)\psi_{,U}^2 \right] .
\]  

(2.25)

Substituting for \( \psi \) from (2.13) and (2.14), and integrating (2.24), (2.25), we obtain

\[
\gamma = \frac{L^2}{4 (\kappa^2 - 1)^4} \left[ \frac{8\kappa^2 + 1}{\rho^4} - \frac{24\kappa(1 + 2\kappa^2)}{\kappa^2 - 1} \frac{U}{\rho^5} + O(\rho^{-6}) \right] \\
+ \frac{J^2}{2 (\kappa^2 - 1)^2} \left[ \frac{1}{\rho^2} - \frac{4\kappa}{\kappa^2 - 1} \frac{U}{\rho^3} + O(\rho^{-4}) \right] .
\]  

(2.26)

Note that we set the integration constant equal to zero. This is because we can go to \( i^+ \) along the center \( \rho = 0 \). More precisely, since we required the regularity of the solution at \( \rho = 0 \), we have to set \( \gamma = 0 \) there and, as a consequence of the field equations for \( \gamma \) in \( (t, \rho) \)-coordinates (cf. (2.8)), \( \gamma \) at \( \rho = 0 \) cannot change with time.

By techniques developed e.g. in [5] it can be now shown that the space-time has a smooth time-like infinity. The analyticity of \( \psi \) at time-like infinity shown in the last sub-section and the field equations imply that the metric “rescaled by inversion” is analytic at \( i^+ \). In what follows, however, we will not use this result; the fall-off properties (2.15) and (2.26) of \( \psi \) and \( \gamma \) will suffice.

### III. NULL INFINITY IN 4 DIMENSIONS

We can now return to the 4-metric (2.1) and analyze its behavior at null infinity. In the main part of this section, we will consider those Einstein-Rosen waves for which the Cauchy data for \( \psi \) in the 3-dimensional picture are (smooth and) of compact support. In the 4-dimensional picture, these solutions correspond to pulses of Einstein-Rosen waves.
A. Formulation of the problem

Let us begin by summarizing the behavior in the directions perpendicular to the axis of symmetry. In these directions, the fall-off of \( \psi \) is the same as in our 3-dimensional treatment of null infinity (see Eq. (2.23) in [1] or Eq. (2.10)). However, from the 4-dimensional perspective, \( \psi \) is not a matter field but a metric coefficient (see (2.1)) and the \( 1/\sqrt{\rho} \) fall-off of \( \psi \) is too slow for null infinity to exist in the sense of Penrose [9]. What is the situation with respect to curvature? In Appendix A, we use the 3-dimensional results to compute the 4-dimensional Riemann tensor for these space-times. We find that, in null directions perpendicular to the \( \partial/\partial z \) Killing field, the tensor decays only as \( 1/\sqrt{\rho} \), the behavior that Stachel first discovered in his direct 4-dimensional treatment [2]. (See the complex components of the Riemann tensor with respect to the null tetrad given by his Eqs. (A4)–(A6), or just \( (4)R_{0030} \), given in our Eq.(A5)). As one would suspect from the behavior of the metric coefficients, the curvature does not peel properly in these null directions. Thus, although we have asymptotic flatness at null infinity of the 3-dimensional space-time [1], the 4-metric fails to be asymptotically flat in null directions perpendicular to the axis of rotation (i.e., along 4-dimensional null lines whose projections approach null infinity in 3 dimensions).

In the rest of the section, we will discuss the fall-off in the remaining null directions. We will find that, contrary to what one might have expected at first, the asymptotic behavior is much better. For time-symmetric initial data, the fall-off in fact satisfies the Bondi-Penrose [9,10] conditions and null infinity is smooth in these directions. Even for a generic data, null infinity exists but may have a “logarithmic behavior”; the conformally rescaled metric is continuous but need not be differentiable there. Note that even this behavior is better than the one in directions orthogonal to the symmetry axis. The reason, in a nutshell, is that the fall-off of various fields along a generic null direction in 4 dimensions is dictated by the fall-off of that field along a time-like direction in the 3-dimensional treatment and, as we saw in Sec. II, fields decay more rapidly at the 3-dimensional time-like infinity than at the 3-dimensional null infinity.

To see this point in detail, let us begin with the Einstein-Rosen metric (cf. (2.1)),

\[
\begin{align*}
ds^2 &= e^{2\psi}dz^2 + e^{2(\gamma-\psi)}(-dt^2 + d\rho^2) + \rho^2 e^{-2\psi}d\phi^2, \\
&= (3.1)
\end{align*}
\]

where \( \psi = \psi(t,\rho) \), \( \gamma = \gamma(t,\rho) \). If we pass from coordinates \( (\rho,z,\phi) \) to spherical coordinates \( (r,\theta,\phi) \), so that \( \rho = r \sin \theta \), \( z = r \cos \theta \), \( \phi = \phi \), and introduce flat–space retarded time \( U = t - r \), we obtain (3.1) in the form

\[
\begin{align*}
ds^2 &= -e^{2(\gamma-\psi)}dU^2 - 2e^{2(\gamma-\psi)}dUd\rho + \left[e^{2\psi} - e^{2(\gamma-\psi)}\right]\cos^2 \theta \, dr^2 \\
&+ \left[e^{2(\gamma-\psi)} \cos^2 \theta + e^{2\psi} \sin^2 \theta\right] \rho^2 d\theta^2 \\
&+ \left[e^{2(\gamma-\psi)} - e^{2\psi}\right] 2r \sin \theta \cos \theta \, dr \, d\theta + r^2 \sin^2 \theta \, e^{-2\psi} \, d\phi^2. \\
&= (3.2)
\end{align*}
\]

Since we are considering waves with initial data of compact support in the \( (\rho,\phi) \)-plane, we can use the results of Section II directly. Recall that one approaches \( i^+ \)
in 3 dimensions, fixing $U = t - \kappa \rho, \kappa = \text{const} > 1$ (cf. (2.11)). In the 4-dimensional picture, this corresponds precisely to approaching the null infinity of the flat metric defined by $t, r, \theta, \phi$ coordinates, along $\theta = \text{const}, \phi = \text{const}, U = t - r = \text{const}$, if we set $\kappa = 1/\sin \theta$. The expansions of $\psi$ and $\gamma$, corresponding to (2.13) and (2.26), thus have the following forms:

$$
\psi = \frac{L}{\cos^3 \theta} \left[ -\frac{1}{r^2} + \frac{(2 + \sin^2 \theta)U}{\cos^2 \theta} \frac{1}{r^3} + O\left(\frac{1}{r^4}\right) \right] + \frac{J}{\cos \theta} \left[ \frac{1}{r} - \frac{U}{\cos^2 \theta} \frac{1}{r^2} + O\left(\frac{1}{r^3}\right) \right],
$$

(3.3)

and

$$
\gamma = \frac{1}{4} \frac{L^2 \sin^2 \theta}{\cos^8 \theta} \left[ (8 + \sin^2 \theta) \frac{1}{r^4} - \frac{24 (2 + \sin^2 \theta)U}{\cos^2 \theta} \frac{1}{r^5} + O\left(\frac{1}{r^6}\right) \right] + \frac{J^2 \sin^2 \theta}{2 \cos^4 \theta} \left[ \frac{1}{r^2} - \frac{4U}{\cos^2 \theta} \frac{1}{r^3} + O\left(\frac{1}{r^4}\right) \right],
$$

(3.4)

provided we stay away from $\theta = \frac{\pi}{2}$, i.e. directions perpendicular to the axis. Our task now is to cast the 4-metric in a Bondi form and show that the metric coefficients have the standard fall-off.

We will carry out this task in the next two sub-sections. We will consider the cases $J = 0, L \neq 0$, and $J \neq 0, L = 0$ separately; since we are interested only in the leading terms, the cross-terms $\sim LJ$ are not relevant. The expansions (3.3), (3.4) show that the fall-off of $\psi$ is slower than that of $\gamma$. Hence, in the construction of the Bondi system, we can focus primarily on $\psi$.

**B. Time-symmetric case; $J = 0$**

$L$ and $J$ are determined by the Cauchy data for scalar waves, and the case $J = 0, L \neq 0$ corresponds to time-symmetric data. Keeping just the first term in $\psi$ in the expansion (3.3) and substituting into (3.2), we find the asymptotic form of the metric to read:

$$
\begin{align*}
    ds^2 &= - \left[ 1 + \frac{2L}{\cos^3 \theta} \frac{1}{r^2} + \ldots \right] dU^2 - 2 \left[ 1 + \frac{2L}{\cos^3 \theta} \frac{1}{r^2} + \ldots \right] dU dr \\
    &- \left[ \frac{4L}{\cos \theta} \frac{1}{r^2} + \ldots \right] dr^2 + \left[ 1 + \frac{2L}{\cos^2 \theta} (\cos^2 \theta - \sin^2 \theta) \frac{1}{r^2} + \ldots \right] r^2 d\theta^2 \\
    &+ \left[ \frac{8L}{\cos^2 \theta} + \ldots \right] \sin \theta dr d\theta + \left[ 1 + \frac{2L}{\cos^3 \theta} \frac{1}{r^2} + \ldots \right] r^2 \sin^2 \theta d\phi^2.
\end{align*}
$$

(3.5)

In order that the metric be of Bondi’s form, we will use the method developed in \cite{11} to analyze space-times with a boost-rotation symmetry. What we need is a coordinate system $\bar{U}, \bar{r}, \bar{\theta}, \bar{\phi} = \phi$ such that

$$
g_{\bar{U}\bar{U}} = 1 + O(\bar{r}^{-1}), \quad g_{\bar{U}\bar{r}} = 1 + O(\bar{r}^{-1})
$$
and, to all orders,
\[ g_{\bar{r}r} = g_{\bar{r}\bar{\theta}} = 0, \quad g_{\bar{\theta}\bar{\theta}} = \bar{r}^4 \sin^2 \bar{\theta}, \]
(3.7)

Let us suppose the transformation leading to this form may be expanded in powers of \( \bar{r}^{-1} \):
\[ U = \pi^0(\bar{U}, \bar{\theta}) + \pi^1(\bar{U}, \bar{\theta})\bar{r}^{-1} + \pi^2(\bar{U}, \bar{\theta})\bar{r}^{-2} + \ldots, \]
\[ r = q(\bar{U}, \bar{\theta})\bar{r} + \sigma^0(\bar{U}, \bar{\theta}) + \sigma^1(\bar{U}, \bar{\theta})\bar{r}^{-1} + \ldots, \]
\[ \theta = \tau^0(\bar{U}, \bar{\theta}) + \tau^1(\bar{U}, \bar{\theta})\bar{r}^{-1} + \tau^2(\bar{U}, \bar{\theta})\bar{r}^{-2} + \ldots. \]
(3.8)

The requirements (3.6), (3.7) restrict the undetermined functions \( \pi, q, \sigma, \tau \). From the leading terms of \( g_{\bar{r}\bar{U}}, g_{\bar{r}\bar{\theta}} \) we first find that \( q_{\bar{U}} = \tau^0_{\bar{U}} = 0 \). The required form of \( g_{\bar{\theta}\bar{\theta}} \) and \( g_{\bar{\theta}\bar{\bar{\theta}}} \) in the leading terms implies \((q)^2\tau^0_{\bar{\theta}} = 1, (q)^2 = \sin^2 \theta / \sin^2 \tau^0 \). This can be solved for \( q \) and \( \tau^0 \) explicitly; however, further we assume \( q = \tau^0_{\bar{\theta}} = 1 \) since the other choices just correspond to coordinate systems connected by boosts along the symmetry axis [10,11]. Then the requirement on the leading order term in \( g_{\bar{U}\bar{r}} \) implies that also \( \pi^0_{\bar{U}} = 1 \). The fall–off conditions (3.6) are thus satisfied.

The conditions \( g_{\bar{\theta}\bar{\bar{\theta}}} = \bar{r}^4 \sin^2 \bar{\theta} + O(\bar{r}^2) \) and \( g_{\bar{r}\bar{\theta}} = O(\bar{r}^{-1}) \) lead to \( \sigma^0 = \tau^1 = 0 \). It remains only to satisfy the requirements (3.7).

The conditions \( g_{\bar{r}\bar{r}} = 0 \) (to order \( O(\bar{r}^{-2}) \)), \( g_{\bar{r}\bar{\theta}} = 0 \) (to \( O(\bar{r}^{-1}) \)) and \( g_{\bar{\theta}\bar{\theta}} g_{\bar{\theta}\bar{\bar{\theta}}} = \bar{r}^4 \sin^2 \bar{\theta} \) (to \( O(\bar{r}^2) \)) determine the functions \( \pi^1, \pi^2, \ldots, \tau^2, \tau^3, \ldots, \) and \( \sigma^1, \sigma^2, \ldots \). More specifically, the vanishing of \( g_{\bar{r}\bar{r}} \) to \( \sim \bar{r}^{-2} \) implies \( \pi^1 = 2L / \cos \bar{\theta}, g_{\bar{r}\bar{\theta}} = 0 \) to \( \sim \bar{r}^{-1} \) leads to \( \tau^2 = -\left(1/2\right)\pi^1 \bar{\theta} + 2L \sin \bar{\theta} / \cos^2 \bar{\theta} \), and \( g_{\bar{\theta}\bar{\bar{\theta}}} = 0 \) in order \( \bar{r}^2 \) gives \( \sigma^1 = -(L / \cos \bar{\theta} + (1/2)\tau^2_{\bar{\theta}}) / \sin \bar{\theta} - (1/2)\tau^2 \cot \bar{\theta} \). To determine the higher order functions \( \pi, \tau \) and \( \sigma \), we have, of course, to consider also the function \( \gamma \) in the metric (3.2). Calculations then become lengthy. Nonetheless, they can be performed and one can thus demonstrate the existence of the Bondi expansion for time-symmetric waves. This establishes the existence of a smooth null infinity in all directions except those perpendicular to the axis of the symmetry.

Now, in axi-symmetric space-times, when a space-like Killing field with circular orbits exists, there is a reduction of the asymptotic symmetry group even if a “global” \( I \) does not exist, i.e., even if \( I \) does not admit spherical cross-sections. Furthermore, in this case, the Bondi news function has a local meaning [12]. One can therefore try to find it in the present case. In Bondi’s coordinates the news function is given by \( c_{\bar{U}} \), where the function \( c(\bar{U}, \bar{\theta}) \) enters, for example, the expansion of \( g_{\bar{\theta}\bar{\bar{\theta}}} = \bar{r}^2 \sin^2 \bar{\theta} + 2c\bar{r} + O(1) \). Starting from our metric (3.2), and using the transformation (3.8) with the functions \( \pi, \sigma, \tau, q \) found above, we obtain \( c = 0 \). Hence the news function vanishes. In fact, this could be anticipated since \( \psi \sim \bar{r}^{-2} \) at \( I \)—we are here in the region in which the tails of cylindrical pulses decay, and there is no radiation field at null infinity [13]. Thus, in these space-times, the radiation
field is all focused in the direction of the two “singular generators” of $I$ singled out by the axis (or, the $\partial/\partial z$-Killing field). Along these generators, the Bondi-Penrose radiation field diverges and asymptotic flatness is lost. In other directions, there is smooth curvature but no flux of energy.

We conclude this sub-section with a remark. In their analysis of isometries compatible with gravitational radiation, Bičák and Schmidt [13] consider axi-symmetric space-times, assume Bondi’s expansion for all $\phi \in [0, 2\pi)$ and $\theta \in (\theta_0, \theta_1)$ and conclude that cylindrical symmetry is not permissible. This assertion may seem to contradict the conclusion we just reached for the time symmetric Einstein-Rosen waves. Note, however, that the interval of permitted $\theta$’s in the assertion of [13] contains $\theta = \frac{\pi}{2}$, i.e., the directions perpendicular to the axis of symmetry, while in the present case, Bondi’s expansion fails to hold in that direction. Thus, there is in fact no contradiction. In fact, the results obtained in the present work are fully compatible with those of [13]; Bičák and Schmidt conclude below their Eq. (52) that, if the function $c = \text{vanishes}$, the second Killing vector field (in addition to the axial one) can generate either a time translation or the translation along the axis of rotation.

**C. Case when $J \neq 0, L = 0$**

In this case, Eq. (3.3) tells us that the leading order behavior of $\psi$ is different: one obtains $\psi \sim J/r \cos \theta$. Consequently, transformation (3.8) does not now lead to a Bondi system; in particular, it does not remove the “offending” term in $g_{rr} \sim r^{-1}$. Nevertheless, since the leading term in the metric does not depend on time and is $O(r^{-1})$, typical for static Weyl metrics, we can attempt to find the required Bondi system by mimicking the procedure adopted in [10]. Let us assume a transformation of the form

$$U = \bar{U} + \pi(r, \bar{\theta}),$$

$$\theta = \bar{\theta} + \tau^1(\bar{\theta}) r^{-1} + \ldots .$$  \hspace{1cm} (3.9)

Keeping then just the first term in $\psi$ in the expansion (3.3) with $L = 0$, and writing the asymptotic form of the metric analogously to (3.5), we find that the crucial term $\sim r^{-1}$ in $\bar{g}_{rr}$ will vanish if

$$- (\pi, r)^2 - 2 \pi, r + \frac{4J}{r} \cos \theta = 0 .$$  \hspace{1cm} (3.10)

Solving in the leading order for $\pi$, we obtain

$$\pi(r, \bar{\theta}) = 2J \cos \bar{\theta} \ln r + \ldots .$$  \hspace{1cm} (3.11)

In this way we can achieve at least $\bar{g}_{rr} \sim O(r^{-2})$. However, with the transformation (3.9) there is no way to satisfy the requirement $\bar{g}_{\bar{\theta} \bar{\theta}} = O(1)$. We must admit a logarithmic term also in the transformation of $\theta$ which, in turn, requires another
logarithmic term in the transformation of \( U \). By assuming expansions in \( r^{-j} \ln^i r \), we find, after some effort, that a suitable transformation reads

\[
U = \bar{U} + \left( \frac{2J}{\cos \theta} \cos^2 \tilde{\theta} \right) \ln r - \left( 2J^2 \sin^2 \tilde{\theta} \right) \frac{\ln^2 r}{r},
\]

\[
\theta = \bar{\theta} + (2J \sin \theta) \frac{\ln r}{r}.
\]

(3.12)

(Notice that in the leading order \( (2J/\cos \theta) \cos^2 \tilde{\theta} = 2J \cos \bar{\theta} \), in agreement with (3.11).)

Now, transforming the metric (3.2), with \( \psi \) and \( \gamma \) given by (3.3), (3.4) (with \( J \neq 0, L = 0 \)), via (3.12), we obtain the metric in the following form:

\[
d\tilde{s}^2 = -\left[ 1 - \frac{2J}{\cos \theta} \frac{1}{r} + O \left( \frac{\ln r}{r^2} \right) \right] d\bar{U}^2
- \left[ 1 - \frac{2J \sin^2 \tilde{\theta}}{\cos \theta} \frac{1}{r} + O \left( \frac{\ln^2 r}{r^2} \right) \right] 2d\bar{U} d\bar{\theta}
+ \left[ 4J \frac{\ln r}{r} + O \left( \frac{\ln^2 r}{r^2} \right) \right] r \sin \bar{\theta} d\bar{U} d\bar{\theta}
+ \left[ O \left( \frac{1}{r^2} \right) \right] dr^2 - \left[ 4J \frac{1}{r} + O \left( \frac{\ln^2 r}{r^2} \right) \right] r \sin \bar{\theta} dr d\bar{\theta}
+ \left[ 1 + O \left( \frac{\ln^2 r}{r} \right) \right] r^2 d\bar{\theta}^2 + \left[ 1 + O \left( \frac{\ln^2 r}{r} \right) \right] r^2 \sin^2 \bar{\theta} d\phi^2.
\]

(3.13)

Bondi et al \cite{10} applied a similar procedure to the Weyl metrics. In contrast to their result, however, we did not quite succeed in bringing our metric to the standard Bondi form. The reason is that, unlike the Weyl metric, in our case, the leading “offending” terms –proportional to \( r^{-1} \)– are \( \theta \)-dependent. \( \text{[In the case of the transformation of the Weyl metric to Bondi’s form –cf. } \cite{10} \text{– we have } \pi = 2m \ln \bar{r} + \ldots, \text{m being the mass. Assuming } m = m(\theta) \text{ in the Weyl metric ( and thus violating the field equations), one can make sure that } I \text{ still exists but the space-time is only “logarithmically” asymptotically flat.] By introducing } \tilde{l} = r^{-1}, \quad \tilde{U} = U, \quad \tilde{\theta} = \theta, \quad \tilde{\phi} = \phi, \text{ and rescaling the metric (3.13) by the conformal factor } \Omega = \tilde{l}, \text{ we obtain}

\[
d\tilde{s}^2 = \Omega^2 d\bar{s}^2 = -\left[ 1 - \frac{2J}{\cos \theta} \tilde{l} + O(\tilde{l}^2 \ln \tilde{l}) \right] \tilde{l}^2 d\bar{l}^2
+ \left[ 1 - \frac{2J \sin^2 \tilde{\theta}}{\cos \theta} \tilde{l} + O(\tilde{l}^2 \ln^2 \tilde{l}) \right] 2d\bar{l} d\tilde{l}
+ \left[ -4J(\sin \tilde{\theta}) \tilde{l} \ln \tilde{l} + O(\tilde{l}^2 \ln^2 \tilde{l}) \right] \tilde{l} d\bar{l} d\tilde{l} + O(1) d\tilde{l}^2
+ \left[ 4J \sin \tilde{\theta} + O(\tilde{l} \ln^2 \tilde{l}) \right] d\tilde{l} d\tilde{\theta}
+ \left[ 1 + O(\tilde{l} \ln^2 \tilde{l}) \right] d\tilde{\theta}^2 + \left[ 1 + O(\tilde{l} \ln^2 \tilde{l}) \right] \sin^2 \tilde{\theta} d\tilde{\phi}^2.
\]

(3.14)
Thus, the metric is well-behaved as \( \tilde{l} \to 0 \), i.e., at \( \tilde{l} = 0 \) \( I \) does exist. The metric is continuous on \( I \). However, it is not differentiable. Thus, it appears that there is a key difference in the asymptotic behavior in the time-symmetric case and in the general case. In the general case, \( I \) appears to have a “logarithmic character” \([3,4]\). (A word of caution is in order: It is possible that the differentiability can be improved by continuing the transformation (3.12) into higher-order terms.)

To conclude, we wish to point out that, although we obtained the asymptotic forms for \( \psi \) and \( \gamma \) (Eqs. (3.3) and (3.4)) assuming that the waves have Cauchy data of compact \( (\rho, \phi) \)-support, the forms themselves hold in more general cases as well. An interesting example is provided by the Weber-Wheeler-Bonnor time-symmetric pulse solution \([14,15]\). [The pulse is formed by a linear superposition of monochromatic waves with a cut-off in the frequency space:

\[
\psi(t,r) = 2C \int_{0}^{\infty} e^{-a\omega} J_0(\omega \rho) \cos \omega t d\omega,
\]

where \( J_0 \) is the Bessel function and the constant \( a \) is an approximate measure of the width of the pulse. It appears \([16]\) that no other integral containing the Bessel function can be expressed in a closed form, which apparently makes the Weber-Wheeler-Bonnor pulse “unique” among non-singular pulse-type solutions of the wave equation in (2+1) dimensions.] In this case, we have

\[
\psi = \sqrt{2C} \left\{ \frac{[(a^2 + \rho^2 - t^2)^2 + 4a^2t^2]^{1/2} + a^2 + \rho^2 - t^2}{(a^2 + \rho^2 - t^2)^2 + 4a^2t^2} \right\}^{1/2}, \quad a = \text{const} \quad (3.15)
\]

and

\[
\gamma = \frac{1}{2} C^2 \left\{ \frac{1}{a^2} - \frac{2\rho^2 \left[ (a^2 + \rho^2 - t^2)^2 - 4a^2t^2 \right]}{[(a^2 + \rho^2 - t^2)^2 + 4a^2t^2]^2} + \frac{1}{a^2 \left[ (a^2 + \rho^2 - t^2)^2 + 4a^2t^2 \right]^{1/2}} \right\}.
\]

(3.16)

At \( t = 0 \), the Cauchy data for \( \psi \) are \( \psi = C(a^2 + \rho^2)^{-\frac{1}{2}} \) and \( \psi_t = 0 \). Nevertheless, expressing the asymptotic forms of \( \psi \) and \( \gamma \) at \( U = t - r = \text{const} \), \( \theta = \text{const} \), \( \phi = \text{const} \), we find after somewhat lengthy calculations (or by using MATHEMATICA), that \( \psi \) and \( \gamma \) have asymptotically exactly the form (3.3), (3.4) with \( J = 0 \) and \( L = -2Ca \). Therefore, in the directions not perpendicular to the symmetry axis, these waves do admit a smooth \( I \).

Similarly, the asymptotic forms of \( \psi \) and \( \gamma \) with \( J \neq 0 \) may hold even though the Cauchy data are not of compact support. A simple prototype, discussed by Carmeli \([17]\), for example, has

\[
\psi = \frac{1}{2\pi} \frac{f_0}{\sqrt{t^2 - \rho^2}}, \quad \gamma = \frac{1}{8\pi^2} \frac{f_0^3 \rho^2}{(t^2 - \rho^2)} , \quad f_0 = \text{const}. \quad (3.17)
\]

This wave is singular at \( t^2 = \rho^2 \) but it represents the late time behavior of the solution given by

\[
\psi = \frac{1}{2\pi} \int_{-\infty}^{\tau} \frac{f(t')dt'}{(t-t')^2 - \rho^2}^{1/2}, \quad \tau = t - \rho , \quad (3.18)
\]
where \( f(t) \neq 0 \) only for \( 0 < t < T \); \( f_0 = \int_0^T f(t') dt' \). With this wave we find \( \psi \) and \( \gamma \) to behave (at \( U = t - r = \text{const}, \theta = \text{const}, \phi = \text{const}, r \to \infty \)) according to (3.3) and (3.4) with \( J = f_0/2\pi \). The fall–off is now slower but a “logarithmic” null infinity, \( I \), does exist.

**Acknowledgements:** AA and JB would like to thank the Albert–Einstein–Institut for its kind hospitality. AA was supported in part by the NSF grants 93-96246 and 95-14240 and by the Eberly Research Fund of Penn State University. JB was supported in part by the grants GACR–202/96/0206 and GA UK–230/1996 of the Czech Republic and the Charles University, and by the US-Czech Science and Technology grant 92067.

**APPENDIX A: APPENDIX: RELATION BETWEEN RIEMANN TENSORS IN 3 AND 4 DIMENSIONS**

The Einstein–Rosen metric (2.1) in coordinates \( x^0 = u = t - \rho, x^1 = \rho, x^2 = \phi, x^3 = z \) becomes

\[
ds^2 = e^{2(\gamma - \psi)}(-du^2 - 2dud\rho) + \rho^2 e^{-2\psi} d\phi^2 + e^{2\psi} dz^2 .
\]

(A1)

Assuming the expansion (2.10) for \( \psi \), we know that \( \gamma \) can be written in the form (2.23), and in principle the Riemann (Weyl) tensor of the vacuum (3+1)–dimensional space-time and its asymptotic behavior can be obtained from (A1).

However, it is possible to use directly the “reduction formulas” for the calculation of the Riemann tensor of spaces which admit an abelian isometry group \([7]\). In the coordinates \((x^\mu) = (u, \rho, \phi, z)\) in which the Killing trajectories orthogonal to the hyper-surfaces \( z = \text{const} \) are just \( u = \text{const}, \rho = \text{const}, \phi = \text{const} \), the 4–dimensional components of the Riemann tensor are given by the following relations (see Eq. (2.3.4) of \([7]\) where, however, the Riemann tensor with the opposite sign is used):

\[
(4) R_{abcd} = \bar{R}_{abcd}, \quad (4) R_{3abc} = 0 ,
\]

\[
(4) R_{3a3b} = -VV_{|ab}, \quad V = e^{\psi} ,
\]

(A2)

where the “\(|\)” denotes the covariant derivative with respect to the metric \( \bar{g}_{ab} \) given by (A1) with \( z = \text{const} \). The covariant derivatives \( V_{|ab} \) are given in terms of the Christoffel symbols listed by Stachel \([2]\) in his Eq. (A1) in Appendix in \((u, \rho, \phi, z)\) coordinates. [In Stachel’s list of \( \Gamma \)’s the following symbols are missing: \( \Gamma_{\rho\phi}^{\rho} = \rho^{-1} - \psi,_{\rho}, \Gamma_{\rho\rho}^{\rho} = 2(\gamma,_{\rho} - \psi,_{\rho}), \Gamma_{\rho z}^{\rho} = \psi,_{\rho} \). Notice also that his \( x^2 = z, x^3 = \phi, \) while here we put \( x^3 = z \). He treats waves with both polarizations so we must put his \( \chi = 0 \) when comparing his results with ours.] Using these,

\[
(4) R_{3030} = -e^{2\psi} \left[ \psi,_{uu} + 3\psi,_{uu}^2 - 2\psi,_{u}\psi,_{\rho} + \psi,_{\rho}^2 + \gamma,_{\rho}(\psi,_{u} - \psi,_{\rho}) + \gamma,_{u}(\psi,_{\rho} - 2\psi,_{u}) \right] ,
\]
\[ (4) R_{3131} = -e^{2\psi} \left[ \psi_{,\rho\rho} + 3\psi^2_{,\rho} - 2\gamma_{,\rho}\psi_{,\rho} \right], \]
\[ (4) R_{3232} = -e^{2\psi - 2\gamma} \rho^2 \left[ 2\psi_{,\rho}\psi_{,\rho} - \psi^2_{,\rho} + \rho^{-1}(\psi_{,\rho} - \psi_{,\rho}) \right], \]
\[ (4) R_{3031} = -e^{2\psi} \left[ \psi_{,\rho\rho} + \psi_{,\rho}\psi_{,\rho} - \gamma_{,\rho}\psi_{,\rho} + \psi^2_{,\rho} \right]. \quad (A3) \]

These are the non-vanishing components \((4) R_{3a3b}\) in the coordinates \((u, \rho, \phi, z)\). Transforming them back to the coordinates \((t, \rho, \phi, z)\) we find — after projecting them on the orthonormal tetrad used by Stachel — precisely his components \( R_{0202}, R_{1212}, R_{2323}\) and \( R_{2021}\) given in his equations \((A3)\). (They have the opposite signs because Stachel uses the signature \(+\ -\ -\ -\).)

The components \( \bar{R}_{abcd}\) (formed from the metric \( \bar{g}_{ab}\)) can be expressed in terms of our \((2+1)\)-dimensional Riemann tensor given in Appendix A of \([1]\). This is formed from the 3–metric \( g_{ab} = e^{2\psi}\bar{g}_{ab}; \) hence we use the behavior of the 3–dimensional Riemann tensor under conformal rescalings (see e.g. \([2]\), Eq. \((2.4.6)\)). We find

\[ \bar{R}_{abjk} = 2e^{2\psi} R_{abjk} - \frac{1}{2} e^{-4\psi}(g_{j[a}V_{b]k} - g_{k[a}V_{b]j}), \quad (A4) \]

where

\[ V_{ik} = 2e^{2\psi} \left[ (g^{lm}\psi_{,l}\psi_{,m})g_{ik} - 2\psi_{,i}\psi_{,k} - 2\psi_{,ik} \right]. \quad (A5) \]

Here the semicolon denotes the covariant derivative with respect to the 3–metric \( g_{ab}\) (see Appendix A of \([1]\) for the Christoffel symbols). The non-vanishing quantities \( V_{ik}\) turn out to be the following:

\[ V_{00} = e^{2\psi} \left[ 2\psi_{,\rho}\psi_{,\rho} - \psi^2_{,\rho} - 2\psi^2_{,u} - 2\psi_{,uu} + 2\psi_{,\rho}(2\gamma_{,\rho} - \gamma_{,u}) + 2\psi_{,\rho}(\gamma_{,\rho} - \gamma_{,u}) \right] \]
\[ V_{01} = V_{10} = 2e^{2\psi} \left[ -\psi^2_{,\rho} - 2\psi_{,u\rho} + 2\gamma_{,\rho}\psi_{,\rho} \right], \]
\[ V_{11} = 4e^{2\psi} \left[ -\psi_{,\rho\rho} - \psi^2_{,\rho} + 2\gamma_{,\rho}\psi_{,\rho} \right], \]
\[ V_{22} = e^{2\psi - 2\gamma} \rho \left( \rho(\psi^2_{,\rho} - 2\psi_{,\rho}\psi_{,\rho}) + 2(\psi_{,u} - \psi_{,\rho}) \right). \quad (A6) \]

By substituting these expressions into \((3.3)\) and using the components \( R_{abcd}\) from Appendix A of \([1]\), we find \( \bar{R}_{abjk} \) — and thus also \((4) R_{abjk}\) — in the coordinates \((u, \rho, \phi, z)\). By transforming them to \((t, \rho, \phi, z)\) we exactly recover Stachel’s expressions given in his Eq. \((A.3)\).
REFERENCES

[1] A. Ashtekar, J. Bičák and B. G. Schmidt, “Asymptotic Structure of Symmetry Reduced General Relativity” (preceding paper).
[2] J. Stachel, J. Math. Phys. 7, 1321 (1966).
[3] J. Winicour, Found. Phys. 15, 605 (1985).
[4] L. Andersson and P. T. Chruściel, Phys. Rev. Lett. 70, 2829 (1993); P. T. Chruściel, M. MacCallum, and D. Singleton, Phil. Trans. R. Soc. London 350, 113 (1995).
[5] J. Bičák and B. G. Schmidt, Phys. Rev. D 40, 1827 (1989).
[6] D. Kramer, H. Stephani, E. Herlt, and M. MacCallum, Exact solutions of Einstein’s field equations (Cambridge University Press, Cambridge, England 1980).
[7] P. Jordan, J. Ehlers, and W. Kundt, Akad. Wiss. Lit. Mainz. Abhandl., Math.–Naturwiss. Klasse, Nr 2 (1960).
[8] G. B. Whitham, Linear and nonlinear waves (J. Wiley, New York 1974).
[9] R. Penrose, Phys. Rev. Lett. 10, 66 (1963); Proc. R. Soc. London A284, 159 (1965).
[10] H. Bondi, M. van der Burg, and A. Metzner, Proc. R. Soc. London A269, 21 (1962).
[11] J. Bičák, Proc. R. Soc. London A302, 201 (1968).
[12] J. Bičák and B. G. Schmidt, Class. Quantum Grav. 6, 1547 (1989).
[13] J. Bičák and B. G. Schmidt, J. Math. Phys. 25, 600 (1984).
[14] J. Weber and J. Wheeler, Rev. Mod. Phys. 29, 509 (1957).
[15] W. B. Bonnor, J. of Math. and Mech. 6, 203 (1957).
[16] I. S. Gradshteyn and I. M. Ryzhik, Table of Integrals, Series, and Products (Academic Press, New York 1980).
[17] M. Carmeli, Classical fields: general relativity and gauge theory (J. Wiley, New York 1982).