THE LAST EIGHT-BILLION YEARS OF INTERGALACTIC Si IV EVOLUTION

KATHY L. COOKSEY1,5, J. XAVIER PROCHASKA2, CHRISTOPHER THOM3, AND HSIAO-WEN CHEN4

1 MIT Kavli Institute for Astrophysics & Space Research, 77 Massachusetts Avenue, 37-611, Cambridge, MA 02139, USA; kcooksey@space.mit.edu
2 Department of Astronomy & UCO/Lick Observatory, University of California, 1156 High Street, Santa Cruz, CA 95064, USA; jxavier@ucolick.org
3 Space Telescope Science Institute, 3700 San Martin Drive, Baltimore, MD 21218, USA; ctom@stsci.edu
4 Department of Astronomy, University of Chicago, 5640 South Ellis Avenue, Chicago, IL 60637, USA; hchen@uchicago.edu

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ABSTRACT

We identified 24 Si iv absorption systems with $z \lesssim 1$ from a blind survey of 49 low-redshift quasars with archival Hubble Space Telescope ultraviolet spectra. We relied solely on the characteristic wavelength separation of the doublet to automatically detect candidates. After visual inspection, we defined a sample of 20 definite (group G = 1) and 4 “highly likely” (G = 2) doublets with rest equivalent widths $W_r$ for both lines detected at $\gtrsim 3\sigma_W$. The absorber line density of the G = 1 doublets was $dN_{Si\text{IV}}/dX = 1.4^{+0.4}_{-0.3}$ for $N(Si^{+3}) > 12.9$. The best-fit power law to the G = 1 frequency distribution of column densities $f(N(Si^{+3}))$ had normalization $k = (1.2^{+0.5}_{-0.4}) \times 10^{-14}$ cm$^2$ and slope $\alpha_N = -1.6^{+0.3}_{-0.3}$. Using the power-law model of $f(N(Si^{+3}))$, we measured the Si$^{+3}$ mass density relative to the critical density: $\Omega_{Si^{+3}} = (3.7^{+2.8}_{-1.7}) \times 10^{-8}$ for $13 \lesssim \log N(Si^{+3}) \lesssim 15$. From Monte Carlo sampling of the distributions, we estimated our value to be a factor of $4.8^{+1.9}_{-1.0}$ higher than the $2 \lesssim z \lesssim 4.5$ ($\Omega_{Si^{+3}}$). From a simple linear fit to $\Omega_{Si^{+3}}$ over the age of the universe, we estimated a slow and steady increase from $z = 5.5 \rightarrow 0$ with $d\Omega_{Si^{+3}}/dt_{age} = (0.61 \pm 0.23) \times 10^{-8}$ Gyr$^{-1}$. We compared our ionic ratios $N(Si^{+3})/N(C^{+3})$ to a $2 < z < 4.5$ sample and concluded, from survival analysis, that the two populations are similar, with median $\langle N(Si^{+3})/N(C^{+3}) \rangle = 0.16$.

Key words: intergalactic medium – quasars: absorption lines – techniques: spectroscopic

Online-only material: color figures, extended figures, machine-readable tables

1. INTRODUCTION

The signatures of the cosmic enrichment cycle are etched into the processed gas and reflected in its metallicity, elemental abundances, density, and/or spatial distribution. Measuring these quantities constrains models of galactic feedback processes (see Bertschinger 1998 and references therein). Currently, quasar absorption-line (QAL) spectroscopy is the best tool for probing the intergalactic medium (IGM).

The ultraviolet (UV) transition Si iv $\lambda\lambda 1393.76, 1402.77$ Å is a well-studied doublet in QAL surveys covering $1.5 < z < 5.5$. From the observability perspective, the Si iv absorption lines are valuable for the following three reasons. First, Si iv absorbers can be observed outside the Ly$\alpha$ forest since they have rest wavelengths redward of Ly$\alpha$ $\lambda 1215$. This reduces the effect of blending. Second, they are observable from ground-based telescopes when they redshift into the optical passband at $z > 1.5$. Third, they constitute a doublet with characteristic rest wavelength separation $(9$ Å) and equivalent width ratio (2 : 1, respectively, in the unsaturated regime). When these distinctive criteria are met, we can be fairly confident that the pair of absorption lines are a Si iv doublet.

From the astrophysics perspective, the Si iv doublet is a strong, observable transition of silicon. The abundance of silicon is predominately driven by Type II supernovae at $z \gtrsim 1$, with an increasing fraction from feedback from asymptotic giant branch stars more recently (Oppenheimer & Davê 2008; Wiersma et al. 2009). Thus, at $z \gtrsim 1$, silicon traces oxygen, the most abundant metal. Oxygen itself is difficult to study since its strong transitions are bluelward of Ly$\alpha$ (e.g., O vi $\lambda\lambda 1031, 1037$), if not also bluelward of the Lyman limit (e.g., O iv $\lambda 787$). By using Si iv absorption as a tracer of oxygen, Songaila (2001) constrained the IGM metallicity to be $> 10^{-5.5}$ Z$_e$ at $z = 5$. Therefore, the fraction of cosmic star formation that occurred before $z = 5$, or within 1 Gyr of the big bang, was $> 10^{-3.5}$.

Songaila (2001) also measured $\Omega_{Si^{+3}}$, the Si$^{+3}$ mass density relative to the critical density, for $2 < z < 5.5$. The Si$^{+3}$ mass density may represent the largest contribution to the silicon mass density $\Omega_{Si}$ at these redshifts (Songaila 2001) but still constitute only a small fraction of $\Omega_{Si}$ (Aguirre et al. 2004). For $z = 4.5 \rightarrow 2$, $\Omega_{Si^{+3}}$ was roughly constant at $\approx 1.2 \times 10^{-8}$ for absorbers with column densities $13 \lesssim \log N(Si^{+3}) \lesssim 15$. For $z = 5.5 \rightarrow 4.5$, the Si$^{+3}$ mass density may have increased by an order of magnitude. Subsequent studies have largely supported these broad trends in $\Omega_{Si^{+3}}$ (non)evolution (Boksenberg et al. 2003; Songaila 2005; Scannapieco et al. 2006).

Observations of gas bearing both Si iv and C iv $\lambda\lambda 1548, 1550$ Å absorption offer constraints on the shape, spatial extent, and/or evolution of the ionizing ultraviolet background (UVB; Boksenberg et al. 2003; Aguirre et al. 2004; Scannapieco et al. 2006). The ionization threshold for Si$^{++}$-to-Si$^{+3}$ is 2.5 Ryd and for C$^{++}$-to-C$^{+3}$, it is 3.5 Ryd. Thus, the ionization ratio $N(Si^{+3})/N(C^{+3})$ is affected by the shape of the UVB at these energies, and the fluctuations in the ratio spatially and in time could constrain the patchiness and evolution, respectively, of the UVB.

Of particular interest is the effect of He ii ionization on the UVB at $z \approx 3$, which is expected to boost Si iv absorption and suppress C iv absorption (Madau & Haardt 2009). Several studies find no evidence for a sharp break in the shape of the UVB at $z \approx 3$ or even significant evolution in its shape for $z \approx 4.5 \rightarrow 1.5$ from studies of Si iv and C iv systems.

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5 NSF Astronomy, and Astrophysics Postdoctoral Fellow.
with column densities of $10^{12} \text{ cm}^{-2}$ to $10^{14} \text{ cm}^{-2}$ (Kim et al. 2002; Boksenberg et al. 2003; Aguirre et al. 2004). Indeed, any variation in the ionic ratio $N(\text{Si}^{14})/N(\text{C}^{+3})$ may be dominated by the variation in the metallicity of the absorbing gas (Bolton & Viel 2010). However, evidence for both a break and strong evolution in the UVB has been detected in some studies of Si$^{14}$ and C$^{+3}$ absorbers (Songaila 1998, 2005).

The gas giving rise to $12 < \log N(\text{Si}^{14}) < 14$ absorption has $T \lesssim 10^{4.9} \text{ K}$ and nearly constant $[\text{Si}/\text{C}] \approx +0.77$ for $1.5 < z < 4.5$ (Aguirre et al. 2004).\footnote{We adopt the following notation: $[\text{Si}/\text{C}] = \log(\sigma_{\text{Si}}/\sigma_{\text{C}}) - \log(\sigma_{\text{Si},0}/\sigma_{\text{C},0})$, where $\sigma_X$ is the volume density of element $X$.} In simulations of the IGM, most (possibly all) of the silicon is located in distinct clouds of metal-enriched gas. At $z \approx 3$, the clouds have radii $\approx 1 \text{ Mpc}$ (proper; Scannapieco et al. 2006) and could be considered filamentary structure. However, by $z = 0$, the enriched clouds are actually the extended gaseous halos ($\approx 100 \text{kpc}$) of galaxies (Davé & Oppenheimer 2007). Absorbers in the circumbinary medium would likely be subjected to a softer ionizing background (due to the increased stellar contribution) than the Haardt & Madau (1996) UVB typically used in IGM studies. Local sources (i.e., star-forming galaxies with a non-zero escape fraction of ionizing photons) may be the most important contributor to the background. If the background for Si$^{14}$ absorbers were softer at $z \lesssim 1$, $[\text{Si}/\text{C}]$ would be lower (Aguirre et al. 2004).

The current work finalizes our analysis of archival Hubble Space Telescope (HST) UV spectra, gathered prior to Servicing Mission 4 (UT 2009 July). This is the largest survey for Si$^{14}$ systems at $z \lesssim 1$ to date and covers the last eight-billion years of the cosmic enrichment cycle. In other words, the net effect of cosmic star formation (and feedback) at the “end” (i.e., $z = 0$) is constrained by observations of Si$^{14}$ absorption at low redshift. Also, the survey of $z \lesssim 1$ Si$^{14}$ absorbers provides a baseline for similar, high-redshift surveys that are currently more numerous. The data reduction and analysis methods used in this paper are described in detail in Cooksey et al. (2010, hereafter Paper I), to which the interested reader is referred.

This paper is organized as follows: we present the data processing and sample selection in Sections 2 and 3, respectively; we analyze and discuss the frequency distribution, its moments, and $N(\text{Si}^{14})/N(\text{C}^{+3})$ in Section 4; and Section 5 is a summary.

2. DATA, REDUCTION, AND MEASUREMENTS

We conducted a blind survey for $z \lesssim 1$ Si$^{14}$ systems in the HST UV spectra of 49 low-redshift quasars, which makes the current work the largest low-redshift Si$^{14}$ study to date. We included spectra from the Space Telescope Imaging Spectrograph (STIS; pre-Servicing Mission 4) and the Goddard High-Resolution Spectrograph (GRHS). The STIS echelle spectra, taken with the E230M grating, provided most of the search path length (see Figure 1), but we also searched the other STIS echelle grating (E140M) and the GHS echelle (ECH-B) and longslit (G160M, G200M, G270M) gratings. Spectra from the Far Ultraviolet Spectroscopic Explorer (FUSE) covered the transitions with rest wavelengths $\lambda_r < 11000 \text{ Å}$ (e.g., higher order H$\alpha$ Lyman lines). All spectra had resolution with full width at half-maximum (FWHM) $\lesssim 15 \text{ km s}^{-1}$ and signal-to-noise ratio $(S/N) \lesssim 2 \text{ pixel}^{-1}$.

The spectra were retrieved from the Multimission Archive at Space Telescope (MAST).\footnote{See http://archive.stsci.edu/.} The reduction and co-addition of multiple observations followed the algorithms described in Cooksey et al. (2008). The spectra were normalized semi-automatically. All reduced, co-added, and normalized spectra are available online, even those not explicitly searched in this paper.\footnote{See http://www.ucolick.org/~xavier/HSTSiIV/ for the normalized spectra, the continuum fits, the Si$^{14}$ candidate lists, and the Monte Carlo completeness limits for all sightlines as well as the completeness test results for the full data sample.}

The central wavelength (and redshift) of an absorption line was measured by the optical depth-weighted mean of the pixels within the wavelength bounds ($\lambda_x, \lambda_b$) defining the absorption line (see Table 1). The rest equivalent widths $W_r$ were measured with simple boxcar summation, and the column densities (e.g., $N(\text{Si}^{14})$) were measured with the apparent optical depth method (AODM; Savage & Sembach 1991). The doublet column density $N(\text{Si}^{14})$ is either: the variance-weighted mean of the measurements of both lines; the column density from the one line with a measurement; the greater lower limit; or the mean, when the limits of the lines define a finite range.

The comoving path length $\Delta X$ sensitive to Si$^{14}$ doublets with $W_r \gtrsim 3 \sigma_{W_r}$ in both lines was estimated from Monte Carlo simulations. For these simulations, we replaced all automatically detected features in the archival spectra with random noise drawn from surrounding pixels. For each redshift bin ($\delta z = 0.005$) of each spectrum, we distributed $10^5$ Si$^{14}$ doublets with a range of column densities, Doppler parameters, and number of components. Then we measured the $N(\text{Si}^{14})$ and $W_r, \lambda_{1393}$ limit per bin at which 95% of the doublets were automatically recovered (see Figure 2). The available path length increases with increasing $N(\text{Si}^{14})$ and $W_r, \lambda_{1393}$; for the strongest absorbers (log $N(\text{Si}^{14}) \gtrsim 14.1$ and $W_{r,1393} \gtrsim 174 \text{ Å}$), $\Delta X = 18$.

3. Si$^{14}$ SAMPLE SELECTION

Absorption lines with observed equivalent width $W_{\text{obs}} \gtrsim 3 \sigma_{W_{\text{obs}}}$ were automatically detected in the spectra. Candidate Si$^{14}$ doublets were identified based solely on the characteristic wavelength separation of the doublet ($\approx 1940 \text{ km s}^{-1}$). First, each $W_{\text{obs}} \gtrsim 3 \sigma_{W_{\text{obs}}}$ line was assumed to be Si$^{14}$ 1393, and...
any automatically detected line that was near the location of the would-be Si iv 1402 was adopted as such. If no $W_{\alpha,\text{obs}} \geq 3\sigma_{W_{\alpha,\text{obs}}}$ line existed, an upper limit was set on $W_{\alpha,\text{obs}}$ from the spectrum. Second, any automatically detected line not already tagged as Si iv 1393 or 1402 was assumed to be the latter, and an upper limit was set on $W_{\alpha,\text{obs}}$ from the spectrum (see Paper I for details). A summary of all candidate Si iv doublets is given in Table 1.

Other common absorption lines (e.g., Ly$\alpha$ and C iv) were associated with the candidate systems in a similar manner. For example, an automatically detected absorption line would be identified as a candidate Ly$\alpha$ line if it had observed wavelength $\lambda \approx \lambda_{\alpha}(1 + z_{\text{cand}})$, where $\lambda_{\alpha}$ is the rest wavelength of Ly$\alpha$ and $z_{\text{cand}}$ is the redshift of the candidate Si iv doublet. The FUSE spectra were useful in this step, since they covered the Ly$\beta$ $\lambda_{1025}$, O vi, and C iii $\lambda 977$ lines for $z \leq 0.1$ candidate systems. The presence of associated lines increased the confidence of our identifications. Each candidate doublet was assigned a machine-generated, binary flag that scaled with the number of desired characteristics of a true Si iv doublet. The characteristics (flags) are as follows:

- 256: $W_{1393} \geq 3\sigma_{W_{1393}}$
- 128: $W_{1402} \geq 3\sigma_{W_{1402}}$
- 64: the ratio $W_{1393} / W_{1402}$ is in the range of 1:1 to 2:1 plus/minus the propagated error of the ratio;
- 32: the optical depth-weighted centroids of the Si iv lines have $|\Delta v_{\text{abs}}| \leq 10$ km s$^{-1}$

10 The wavelength separation between a line at $z_{\text{abs}}$ and a Si iv line at $z_{1393}$ is $\Delta v_{\text{abs}} = c(z_{\text{abs}} - z_{1393})/(1 + z_{1393})$.

16: there exists a candidate Ly$\alpha$ line with $W_{r,\alpha} \geq 3\sigma_{W_{r,\alpha}}$
8: the 1393 line is outside of the Ly$\alpha$ forest
4: and outside the H$\alpha$ forest;
2: the smoothed AOD per pixel of the doublet lines agree within 1σ for $\geq 68.3\%$ of the pixel; and
1: there exists one or more candidate lines (not H1) associated with the candidate doublet and with $W_{r} \geq 3\sigma_{W_{r}}$.

All candidates were visually inspected by at least one author, and the candidates with both lines detected at $W_{r} \geq 3\sigma_{W_{r}}$ were reviewed by two or more. We agreed upon 22 definite Si iv systems, which constitute the “G = 1” group (see Figure 3). Of these, 20 have both doublet lines detected with rest equivalent width $W_{r} \geq 3\sigma_{W_{r}}$ and constitute the group on which we based our conclusions. The G = 1 sample has a median redshift $(z) = 0.906$, log $N$(Si iv) $> 12.9$, and $W_{r,1393} \geq 66$ Å.

We also defined a small, “highly likely” (G = 2) sample of six systems (see Figure 4). These doublet systems are typically found in regions with low S/N and/or do not have other lines associated with them, which would increase the confidence of our identification. The four of these with both lines detected at $W_{r} \geq 3\sigma_{W_{r}}$ were combined with the G = 1 sample for some analyses and identified in tables and figures as G = 1+2. Details of all (G = 1+2) absorption systems are given in Table 2. The properties of all the Si iv doublets are summarized in Table 3.

All doublets are more than 1000 km s$^{-1}$ outside of the Galaxy and 3000 km s$^{-1}$ blueward of the background quasar. There were no Si iv absorbers without associated Ly$\alpha$ absorption, when the spectral coverage existed to detect Ly$\alpha$. We combined Si iv
doublets into one absorption system when their optical depth-weighted centroids had $\delta v_{\text{abs}} < 250 \text{ km s}^{-1}$.

The observed absorber line density is the sum of the number of absorbers, each weighted by the path length sensitive to their $N(\text{Si}^{+3})$ or $W_r(1393)$. For the $G = 1$ sample, $dN_{\text{Si}^{+3}}/dX = 1.4_{-0.3}^{+0.4}$ ($dN_{\text{Si}^{+4}}/dz = 2.7_{-0.6}^{+0.7}$) for $\log N(\text{Si}^{+3}) > 12.9$.

As in Paper I, we conducted Monte Carlo simulations to measure the rate that pairs of $z \lesssim 1.5 \text{ H}I$ Lyman forest lines satisfy the characteristics of $\text{Si}^{+4}$ doublets and were potentially included in our sample as doublets. The contamination rate of forest lines masquerading as $\text{Si}^{+4}$ was small, less than 5% of $dN_{\text{Si}^{+4}}/dX$ or an expected $1_{-0.7}^{+1}$ false doublet. If any doublet in our sample were false, it would be one without other associated absorption lines. There are four such doublets, all in the $G = 2$ sample. Our expectation that forest lines could mimic $\text{Si}^{+4}$ doublets drove us to define the “highly likely” $G = 2$ group. Though we provide the results for analyses of the $G = 1+2$ sample, we only discuss the results from the $G = 1$ sample and base our conclusions on that; so, we concern ourselves no further with the effects of the $\text{H}I$ Lyman forest contamination.

### 3.1. Comparison with Previous Studies

There have been three recent surveys for $\text{Si}^{+4}$ systems at $z \lesssim 1$ using at least some of the HST spectra analyzed here: Milutinović et al. (2007), Danforth & Shull (2008), and Paper I. Here we briefly compare our blind doublet search results with these other studies. In the first two, they identified Ly$\alpha$ lines first and then sought associated transitions.

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**Figure 2.** Redshift path length $\Delta X(\log N(\text{Si}^{+3}))$ and $\Delta X(W_r(1393))$ as a function of $\text{Si}^{+4}$ column density (top) and rest equivalent width (bottom), respectively. These estimates are based on Monte Carlo analysis and correspond to 95% completeness limits.

**Figure 3.** Velocity plots of 22 $G = 1 \text{ Si}^{+4}$ systems. The regions of spectra around each absorption line are aligned in velocity space with respect to the rest wavelength of the transition and $z(1393)$. Saturated transitions are indicated with the (red) “S”; any transition with column density detected at <3$\sigma$ are noted with a (red) “W.” Systems with $z(1393) = \lambda_{\alpha}(1 + z_{\text{abs}})$ are labeled with a (red) “Ly$\alpha$ Forest.” The regions used to measure $W_r$ and $\log N(\text{Si}^{+3})$ are shown by the dark outline. The flux at zero and unity are shown with the dash-dotted lines (blue and green, respectively); the vertical dashed line indicates $v = 0 \text{ km s}^{-1}$, corresponding to the optical depth-weighted velocity centroid of the $\text{Si}^{+4}$ 1393 transition. (The $G = 1$ velocity plots are available in their entirety in the online journal. One figure is shown here for guidance regarding the form and content.) (A color version and an extended version of this figure are available in the online journal.)

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**Figure 4.** Velocity plots of six $G = 2 \text{ Si}^{+4}$ systems. (See Figure 3 for description of velocity plot.) (The $G = 2$ velocity plots are available in their entirety in the online journal. One figure is shown here for guidance regarding the form and content.) (A color version and an extended version of this figure are available in the online journal.)
were explicitly listed as questionable by these authors. We identified 13 of those absorbers. The remaining four doublets are labeled group G and were doublets from Danforth & Shull (2008) were not included in our sample for at least one of the following reasons: one or both doublet lines were at z\em \lambda_r \lambda_l \lambda_h \sigma_W \log N \sigma_{\log N} G Flag
PG0117+213 (z_{em} = 1.493)
0.65536 ... 1393.76 2306.70 2307.89 273 28 >13.73 ... 2 482
9.2 1402.77 2321.62 2322.82 122 24 13.54 0.09
1.04801 ... 1393.76 2853.75 2854.89 194 18 13.54 0.07 1 471
−10.5 1402.77 2872.21 2873.36 83 14 13.37 0.07
−111.3 1215.67 2486.96 2490.84 1530 21 >14.88 ... 2
−27.2 1206.50 2469.68 2471.49 446 19 >13.67 ... 2
−26.2 1260.42 2580.19 2581.61 145 11 13.16 0.03
B0312−770 (z_{em} = 0.223)
0.20253 ... 1393.76 1674.90 1676.82 652 78 >14.16 ... 1 511
7.6 1402.77 1685.73 1687.67 543 57 >14.40 ...
−15.0 1215.67 1460.66 1463.09 1677 20 >14.92 ...
−7.3 1025.72 1232.52 1234.29 1139 20 >15.59 ...
11.2 977.02 1174.06 1175.78 912 67 >14.44 ...
51.6 1031.93 1240.64 1241.62 496 17 >14.94 ...
58.2 1037.62 1247.48 1248.07 310 14 14.87 0.02
21.7 1206.50 1449.82 1451.60 763 19 >13.94 ...
61.5 1238.82 1489.57 1490.45 137 7 13.86 0.02
62.1 1242.80 1494.36 1495.24 94 9 13.99 0.04
39.1 1260.42 1515.44 1516.45 490 12 >13.92 ...
PK03405−123 (z_{em} = 0.573)
0.16712 ... 1393.76 1626.40 1626.95 117 15 13.24 0.06 1 499
5.4 1402.77 1636.92 1637.47 103 23 13.49 0.10
−18.1 1215.67 1417.80 1419.38 876 19 >14.60 ...
−0.9 1025.72 1196.77 1197.50 444 20 >15.19 ...
−22.2 977.02 1139.53 1140.73 500 22 >14.26 ...
12.1 989.80 1154.85 1155.66 214 12 >14.64 ...
−30.1 1031.93 1203.74 1204.67 391 24 >14.77 ...
−50.7 1037.62 1210.37 1211.31 218 37 14.80 0.11
2.3 1206.50 1407.81 1408.40 236 11 >13.38 ...
−6.4 1238.82 1445.53 1446.12 108 15 13.81 0.06
−7.4 1242.80 1450.17 1450.76 76 13 13.94 0.07
−3.3 1260.42 1470.73 1471.28 148 12 >13.28 ...
Notes. Siiv systems by target and redshift of Siiv 1393. Upper limits are 2$\sigma$ limits for both W_r and log N. The column densities were measured by the AODM, unless “COG” is indicated, in which case, the limit is from assuming W_r results from the linear portion of the COG. The definite Siiv doublets are labeled group G = 1, while the “highly likely” doublets are G = 2. The binary flag is described in Section 3.
(This table is available in its entirety in a machine-readable form in the online journal. A portion is shown here for guidance regarding its form and content.)

Milutinović et al. (2007) identified 17 Siiv doublets in the eight STIS E230M spectra that they surveyed. We independently identified 13 of those absorbers. The remaining four doublets were explicitly listed as questionable by these authors. We identified the G = 2, z_{1393} = 0.65536 absorber in the PG0117+213 spectrum, one of the eight surveyed by Milutinović et al. (2007), though they did not detect it. Since only the doublet was detected, they would not have found it with their Ly$\alpha$-targeted search.

Danforth & Shull (2008) surveyed all of the STIS E140M spectra for many transitions, including Siiv. They did not require that both doublet lines be detected with W_r \geq 3$\sigma$W_r. They found 20 Siiv doublets for a line density dN_{Siiv}/dX = 5$^{15}$ for W_r \geq 30 Å.

In contrast, we found seven Siiv absorbers in the E140M spectra. There were two (G = 1) doublets that we identified in our Siiv-targeted survey but Danforth & Shull (2008) did not find: the z_{1393} = 0.00572 doublet in QSO–123050+011522 and the z_{1393} = 0.13846 one in PG1116+215. As mentioned in Paper I, Danforth & Shull (2008) missed the former doublet because the Ly$\beta$ line in the FUSE spectra was suspect. The latter doublet has W_r,1402 < 3$\sigma$W_r,1402, which might explain why they did not identify it.

Of their 20 systems, we agree with five and independently identified them in our Siiv-targeted survey. The remaining 15 doublets from Danforth & Shull (2008) were not included in our sample for at least one of the following reasons: one or both lines detected were at W_r < 3$\sigma$W_r (12 doublets); one line was blended with a Galactic line (1); it was a system intrinsic to the background quasar (1); and/or the doublet was observed in the zeroth order of E140M (2), which was excluded in our reduction.

We recovered all Siiv doublets that we previously identified in Paper I because of the association with a C iv system. We also determined that the G = 1, z_{1548} = 0.24010 C iv doublet from Paper I is actually O i λ1302 and Siiv λ1304 associated with the G = 1, z_{1548} = 0.47436 system, which corresponds to our G = 1 Siiv system with z_{1393} = z_{1548}. 
4. ANALYSIS

4.1. Frequency Distributions

Analogous to the luminosity function used in galaxy studies, observers define the column density frequency distribution \( f(N(\text{Si}^+)) \) to be the number \( \Delta N \) of \( \text{Si}^+ \) absorbers per column density interval \( \Delta N(\text{Si}^+) \) per path length \( \Delta X \) (see Figure 5). With a maximum likelihood fitting algorithm, we fit a power law to \( f(N(\text{Si}^+)) \) as follows:

\[
f(N(\text{Si}^+)) = k \left( \frac{N(\text{Si}^+)}{N_0} \right)^{\alpha_N},
\]

where \( k \) is the normalization with unit \( \text{cm}^{-2} \), \( N_0 = 10^{13.5} \text{ cm}^{-2} \), and \( \alpha_N \) is the slope. We often refer to \( k_{14} \), which is the normalization \( k \) scaled up by a factor of 10\(^{14} \). The frequency distribution was fitted over the range \( \log N_{\text{min}} = 12.84 \) to \( \log N_{\text{max}} = 15 \), with special treatment of the saturated absorbers (\( \log N_{\text{sat}} = 14 \)) in the maximum likelihood analysis (see Paper I). Briefly, the number of doubles with \( 14 < \log N(\text{Si}^+) < 15 \) (i.e., saturated) was a constraint in our likelihood function. The best-fit parameters were \( k_{14} = 1.18^{+0.45}_{-0.36} \text{ cm}^{-2} \) and \( \alpha_N = -1.61^{+0.28}_{-0.31} \).

The fitted parameters for \( f(N(\text{Si}^+)) \) and \( f(W_{\text{e,1393}}) \) (discussed below) are given in Table 4. We estimated the 68.3% confidence limits (c.l.) in the power-law normalization and slope by tracing a contour where \( \delta L \approx \ln L - \ln L_{\text{max}} = -1.15 \) on the likelihood surface \( L \), which included 68.3% of its area. Then, the 68.3% c.l. (what we will loosely refer to as “1σ errors,” hereafter) were defined as the difference between the \( k \) and \( \alpha \) extrema on the contour and the \( L_{\text{max}} \) values. Errors in quantities derived from the frequency distributions (e.g., the \( \text{Si}^+ \) mass density discussed below) are estimated in a similar fashion. The likelihood surface is not Gaussian, and the “2σ and 3σ errors” are defined by \( \delta L = -3.15 \) (95.4% c.l.) and \(-6.1\) (99.7% c.l.), respectively. Therefore, the larger confidence limits on the best-fit power-law parameters are, formally, as follows:

\[
\begin{align*}
\delta \alpha &= +0.79/ -0.54 \text{ cm}^2; & \delta k &= +0.45/ -0.53; \\
3\delta \alpha &= +1.18/ -0.70 \text{ cm}^2; & 3\delta k &= +0.62/ -0.76.
\end{align*}
\]

There was no observed break in \( f(N(\text{Si}^+)) \), and no break has been observed at high redshift. There must be a break in order to limit the number and mass of \( \text{Si}^+ \) absorbers to finite quantities.

We have measured a slope consistent with those from high-redshift studies, though potentially shallower. Songaila (1997) measured \( \alpha_N = -1.8 \) for \( 2.16 \leq z \leq 3 \) and \(-2 \) for \( 3 \leq z \leq 3.54 \). Songaila (2001) and Scannapieco et al. (2006) stated that \( \alpha_N = -1.8 \) matched their observed \( f(N(\text{Si}^+)) \) well, which covered \( 1.78 \leq z \leq 5.29 \) and \( 1.5 \leq z \leq 3.1 \), respectively. From their Ly\( \alpha \)-targeted \( z < 0.4 \) survey, Danforth...
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Table 4
Si iv Frequency Distributions Summary

| (1) G | (2) z | (3) ziv | (4) N | (5) Limits | (6) dN_{Siiv}/dz | (7) dN_{Siiv}/dX | (8) \( \Omega_{Siiv} \times 10^4 \) | (9) k | (10) \( \alpha \) | (11) PKS |
|------|------|--------|------|-----------|-----------------|-----------------|-----------------|-----|-------|--------|
| 1    | 0.90560 | 0.00428 | 1.14760 | 20 (12.84, 15.00) | 2.71^{+1.1}_{-0.8} | 1.2^{+0.5}_{-0.4} | 3.71^{+2.8}_{-2.2} | -1.66 | 1.18^{+0.45}_{-0.36} | -1.61^{+0.28}_{-0.31} | 0.526 |
| 0.91432 | 18 (12.92, 14.40) | 2.4^{+0.4}_{-0.3} | 1.4^{+0.2}_{-0.2} | >1.63 | \ldots | \ldots | \ldots | \ldots | \ldots | \ldots | \ldots |
| 1+2  | 0.85485 | 0.00428 | 1.14760 | 24 (12.84, 15.00) | 3.1^{+1.1}_{-1.0} | 1.5^{+0.5}_{-0.4} | 4.13^{+2.81}_{-1.77} | 1.43^{+0.48}_{-0.39} | -1.65^{+0.26}_{-0.28} | 0.400 |
| 0.90560 | 22 (12.92, 14.40) | 2.9^{+0.4}_{-0.3} | 1.7^{+0.2}_{-0.2} | >1.94 | \ldots | \ldots | \ldots | \ldots | \ldots | \ldots | \ldots |

Notes. Parameters from the maximum likelihood analysis for \( f(x) = k(x/x_0)^\alpha \), where \( x = N(Si iv) \) (or \( W_{r,1393} \)) and \( x_0 \) is \( N_0 = 10^{13.5} \text{cm}^{-2} \) (\( W_{r,0} = 150 \text{Å} \)). For each \( G = 1 \) or 1+2 subsample, the first row summarizes the maximum likelihood analysis and the second row, the observed quantities, \( dN_{Siiv}/dX \), listed in the first subsample row, is the integral of \( (N(Si iv)) \) \( f(W_{r,1393})) \) from 13 to infinity (\( W_{r,1393} = 50 \text{Å} \) to 796 Å) with the best-fit \( k \) and \( \alpha \). Also, in the first subsample row, the integrated \( dN_{Siiv}/dX = dN_{Siiv}/dX \cdot dX/dz \), where the latter term is evaluated at \( z \). The observed \( dN_{Siiv}/dX \) and \( dN_{Siiv}/dX \) are from the sum of the total number of doublets, weighted by the path length available to detect the doublet, based on the observed \( N(Si iv) \) \( W_{r,1393} \) at \( \Omega_{Siiv} \), listed in the first subsample row in the column density section, is the integral of \( f(N(Si iv)) - N(Si iv) \) from 13 \( \log N(Si iv) \) \( <15 \) with the best-fit \( k \) and \( \alpha \). The observed \( \Omega_{Siiv} \), were from the sum of the unsaturated doublets, as given by \( N \). The significance of the one-sided Kolmogorov–Smirnov statistic of the best-fit power law.

\( k \) and \( \alpha \) of \( (Si iv) \) measured \( \alpha_N = -1.92 \pm 0.17 \) for a sample with different selection criteria and from a survey with different methodology.

The definition of the equivalent width frequency distribution \( f(W_{r,1393}) \) is similar to that of \( f(N(Si iv)) \), and it was also fitted well with a power law, with \( W_{r,0} = 150 \text{Å} \), \( W_{r,1393} = 57 \text{Å} \), and \( W_{r,1393} = 796 \text{Å} \). The \( W_r \) limits reflect the extrema of the observed values, \( \pm 1\sigma \). The best-fit values for \( f(W_{r,1393}) \) were \( k_5 = 3.48^{+1.4}_{-1.12} \) and \( \alpha_W = -1.28^{+0.47}_{-0.47} \). The \( W_r \) are the normalization scaled up by a factor of 10^3. The larger confidence limits on the best-fit power-law parameters for \( f(W_{r,1393}) \) are, formally, as follows: \( 2\sigma_k = +2.55/1.73 \text{Å}^{-1} \), \( 2\sigma_{\alpha_W} = +0.79/0.80 \), and \( 3\sigma_k = +3.81/2.21 \text{Å}^{-1} \), and \( 3\sigma_{\alpha_W} = +1.10/1.13 \).

4.2. Si iv Absorber Line Density

We measured the observed Si iv line density by the sum of the number of absorbers, each weighted by the path length sensitive to their \( N(Si iv) \) or \( W_{r,1393} \). We found that \( \int f(W_{r,1393}) \) was well modeled by a power law, we can integrate Equation (1) to estimate the Si iv absorber line density \( dN_{Siiv}/dX \) for a given column density limit \( N_{lim} \):

\[
\frac{dN_{Siiv}}{dX} (N(Si iv) \geq N_{lim}) = \frac{-1}{k + \alpha_N} N_{lim}^{1+\alpha_N} N_0^{-\alpha_N}. \tag{2}
\]

This is useful for comparing to high-redshift studies, which typically do not match our observational limits. In Figure 6, we show \( dN_{Siiv}/dX \) as a function of \( z_{1393} \) for \( N_{lim} = 10^{13} \text{cm}^{-2} \).

The Si iv line density does not increase significantly from \( z \approx 3 \to 0 \). Our integrated, \( z \lesssim 1 \frac{dN_{Siiv}}{dX} = 1.2^{+0.5}_{-0.4} \).
The formal errors on the integrated $\rho_c$, $f$ and $k$ (log $N$ of the universe sample). Doublets, which dominate the critical density $\rho_c$ at several $z$. We excluded the two saturated doublets since they might have their $\rho_c$ estimate of the errors. We assumed the $\rho_c$ values for $N(Si^{+3})$ published in Scannapieco et al. (2006, black squares). (A color version of this figure is available in the online journal.)

$log N_{lim} = 13$ is consistent within $1\sigma$ of the $\langle z \rangle = 1.9$ value from Scannapieco et al. (2006) and $2\sigma$ of their $\langle z \rangle = 2.7$ one. The formal errors on the integrated $dN_{SiIV}/dX$ are $2\sigma_{dN/dX} = +1.3/-0.6$ and $3\sigma_{dN/dX} = +1.2/-0.8$, derived from the $\alpha_N$ and $k_{14}$ errors discussed in Section 4.1. There is no evidence of evolution in $dN_{SiIV}/dX$ from $z = 1 \rightarrow 0$ (i.e., within our sample).

We estimated the high-redshift line densities and errors in the following manner. Scannapieco et al. (2006) published their $f(N(Si^{+3}))$, from which we estimated the power-law normalization $k_{14}$ given the best-fit slope $\alpha_N$. We measured $f(N(Si^{+3}))$ at several $N(Si^{+3})$ and used the scatter as an estimate of the errors. We assumed $\alpha_N = -1.8$, as Scannapieco et al. (2006) did, and estimated $k_{14} = 1.3 \pm 0.9$ cm$^{-2}$ for $\langle z \rangle = 1.9$ and $k_{14} = 0.5 \pm 0.1$ cm$^{-2}$ for $\langle z \rangle = 2.7$.

### 4.3. Si$^{+3}$ Mass Density

The Si$^{+3}$ mass density has been measured by several $1.5 < z < 5.5$ studies (Songaila 1997, 2001; Scannapieco et al. 2006). Typically, they measured the mass density, relative to the critical density $\rho_c$, by summing the detected absorbers (Lanzetta et al. 1991):

$$\Omega_{Si^{+3}} = \frac{H_0 m_{Si}}{c \rho_c,0} \sum N(Si^{+3}) \Delta X (N(Si^{+3})), \quad (3)$$

where the Hubble constant is $H_0 = 70$ km s$^{-1}$ Mpc$^{-1}$, the mass of a silicon atom is $m_{Si} = 5 \times 10^{-23}$ g, $c$ is the speed of light, and $\rho_c,0 = 9.26 \times 10^{-30}$ g cm$^{-3}$. The summed $\Omega_{Si^{+3}}$ approximates the mass density in all absorbers with column densities within the observed $N(Si^{+3})$ range, which is $12 \lesssim log N(Si^{+3}) \lesssim 15$ for the high-redshift studies.

Since we measured column densities with the AODM, we only have lower limits on $N(Si^{+3})$ for the strong, saturated doublets, which dominate $\Omega_{Si^{+3}}$. Thus, summing the column densities only resulted in a lower limit $\Omega_{Si^{+3}} \gtrsim 1.55 \times 10^{-8}$ for the 16 unsaturated G = 1 doublets with $log N(Si^{+3}) \gtrsim 13$. We exclude the two saturated doublets since they might have $log N(Si^{+3}) > 15$, which is the upper limit we have chosen for comparing $\Omega_{Si^{+3}}$. Including the two saturated doublets results in the summed $\Omega_{Si^{+3}} > 2.3 \times 10^{-8}$ for the 18 G = 1 doublets with $log N(Si^{+3}) \gtrsim 13$.

In order to compare with the high-redshift studies, we assumed the power-law formalism for $f(N(Si^{+3}))$ and integrated the column density “weighted” $f(N(Si^{+3}))$ (i.e., its first moment):

$$\Omega_{Si^{+3}} = \frac{H_0 m_{Si}}{c \rho_c,0} \frac{k}{2 + \sigma_N} \left( \frac{N^{2+\alpha_N} - N^{2+\alpha_N}_{min}}{N^{2+\alpha_N}_{max} - N^{2+\alpha_N}_{min}} \right). \quad (4)$$

From the best-fit values for the $G = 1$ $f(N(Si^{+3}))$, $\Omega_{Si^{+3}} = (3.71^{+2.82}_{-1.68}) \times 10^{-8}$ for $log N_{min} = 13$ and $log N_{max} = 15$. Since there has been no observed break in $f(N(Si^{+3}))$, the column density limits are crucial to defining a finite $\Omega_{Si^{+3}}$ and comparing between surveys.

The formal errors on the integrated $\Omega_{Si^{+3}}$ are $2\sigma_{\Omega} = +5.4/-2.36$ and $3\sigma_{\Omega} = +8.99/-2.81$, derived from the $\alpha_N$ and $k_{14}$ errors discussed in Section 4.1.

We plot the evolution of $\Omega_{Si^{+3}}$ as a function of the age of the universe $t_{age}$ in Figure 7. The median $11$ of the $2 \lesssim z < 4.5$ studies (Songaila 2001; Scannapieco et al. 2006, whose values have been adjusted to match our $N(Si^{+3})$ limits and cosmology, see Paper I) is $\langle \Omega_{MC}^{Si^{+3}} \rangle (z < 4.5) = (0.77^{+0.31}_{-0.21}) \times 10^{-8}$, which is shown by the (blue) lines in Figure 7 (left). We have detected, with $>99.8\%$ confidence, an increase in $\Omega_{Si^{+3}}$ from high to low redshift. The $z \lesssim 1$ is a factor of $4.8^{+3.0}_{-1.9}$ higher than the $2 \lesssim z \lesssim 4.5$ median.

The confidence limits on the median high-redshift value and the increase in $\Omega_{Si^{+3}}$ were estimated based on Monte Carlo sampling of the distributions. First, we drew $10^4$ realizations of the high-redshift data sets $\Omega_{MC}^{Si^{+3}}$, assuming they had log-normal errors. Then, we measured the median of each set: $\langle \Omega_{MC}^{Si^{+3}} \rangle$. The median Si$^{+3}$ mass density quoted above ($\langle \Omega_{MC}^{Si^{+3}} \rangle$) was the median of these ($10^4$) $\langle \Omega_{MC}^{Si^{+3}} \rangle$ values.

To Monte Carlo sample our $z \lesssim 1$ measurement, we used the likelihood surface discussed in Section 4.1 and the Metropolis–Hastings algorithm to appropriately sample the $k$–$\alpha_N$ parameter space $10^4$ times. For each of these random pairs, we computed the integrated Si$^{+3}$ mass density (see Equation (4)), resulting in a low-redshift Monte Carlo sample: $\Omega_{MC}^{Si^{+3}}$. The median of this distribution was $\langle \Omega_{MC}^{Si^{+3}} \rangle = (3.70^{+1.68}_{-1.11}) \times 10^{-8}$.

The ratio of the low- to high-redshift Monte Carlo samples (i.e., $\langle \Omega_{MC}^{Si^{+3}} \rangle / \Omega_{MC}^{Si^{+3}}$) is a distribution where the median is $4.8^{+3.0}_{-1.9}$ and the ratio is greater than unity at the $99.8\%$ c.l.

A least-squares minimization of a linear model to $\Omega_{Si^{+3}}$ over $t_{age}$ for the $2 < z < 5.5$ observations (Songaila 2001; Scannapieco et al. 2006) and our $z \lesssim 1$ value yielded: $d\Omega_{Si^{+3}}/dt_{age} = (0.61 \pm 0.23) \times 10^{-8}$ Gyr, as shown by the (red) lines in Figure 7 (right). This toy-model slope agrees well ($<2\sigma$) with $d\Omega_{MC}^{Si^{+3}}/dt_{age} = (0.42 \pm 0.2) \times 10^{-8}$ Gyr from Paper I for the equivalent redshift sample but for CIV absorbers, and D’Odorico et al. (2010) detected a smooth increase in $\Omega_{Si^{+3}}$ from $z = 2.5 \rightarrow 1.5$, which reaches the $z < 1$ $\Omega_{Si^{+3}}$ limit.

11 When we quote median values with “1σ errors,” these “errors” are actually the difference between the median (i.e., 50th percentile) and the values at the 15.9th and 84.2nd percentiles.

12 This value is in excellent agreement with our integrated value for the best-fit $k$ and $\alpha_N$ and indicates a smaller spread, within the quoted confidence limits, than we formally adopted (see Table 4).
measured in Paper I. A linear fit to only the $2 \leq z < 5.5$ observations indicated no temporal evolution in $\Omega_{\text{Si}^{\text{i+3}}}$, for $13 \leq \log N(\text{Si}^{\text{i+3}}) \leq 15$ for the G = 1 sample, is the (orange) star. The value from Danforth & Shull (2008) is the (green) triangle and is not an independent measurement of $\Omega_{\text{Si}^{\text{i+3}}}$. We compared our $\Omega_{\text{Si}^{\text{i+3}}}$ measurement only to those from Songaila (2001, black crosses) and Scannapieco et al. (2006, black, open square). The median, and its “errors” (see Section 4.3), of the $2 \leq z \leq 4.5$ values are shown on the left with the (blue) solid and dashed lines, respectively. The simple linear fit to $\Omega_{\text{Si}^{\text{i+3}}}$ over $\log_{10} z_{\odot}$ is shown on the right by the (red) solid line, where the dashed lines are the 1σ range of the fit. For reference, we also show the values from Songaila (1997, gray circles) and Songaila (2005, gray, filled squares), the latter of which uses a pixel-optical depth method to measure $\Omega_{\text{Si}^{\text{i+3}}}$.

(A color version of this figure is available in the online journal.)

Figure 7. Time evolution of $\text{Si}^{\text{i+3}}$ mass density relative to the critical density. The integrated $\Omega_{\text{Si}^{\text{i+3}}}$, for $13 \leq \log N(\text{Si}^{\text{i+3}}) \leq 15$ for the G = 1 sample, is the (orange) star. The value from Danforth & Shull (2008) is the (green) triangle and is not an independent measurement of $\Omega_{\text{Si}^{\text{i+3}}}$. We compared our $\Omega_{\text{Si}^{\text{i+3}}}$ measurement only to those from Songaila (2001, black crosses) and Scannapieco et al. (2006, black, open square). The median, and its “errors” (see Section 4.3), of the $2 \leq z \leq 4.5$ values are shown on the left with the (blue) solid and dashed lines, respectively. The simple linear fit to $\Omega_{\text{Si}^{\text{i+3}}}$ over $\log_{10} z_{\odot}$ is shown on the right by the (red) solid line, where the dashed lines are the 1σ range of the fit. For reference, we also show the values from Songaila (1997, gray circles) and Songaila (2005, gray, filled squares), the latter of which uses a pixel-optical depth method to measure $\Omega_{\text{Si}^{\text{i+3}}}$. (A color version of this figure is available in the online journal.)

A linear fit to only the $2 \leq z < 5.5$ observations indicated no temporal evolution in $\Omega_{\text{Si}^{\text{i+3}}}$ (i.e., $d\Omega_{\text{Si}^{\text{i+3}}}/dt_{\text{age}} = (0.10 \pm 0.56) \times 10^{-8}$ Gyr).

Modeling $\Omega_{\text{Si}^{\text{i+3}}}$ as evolving linearly with time was not a physically motivated exercise but one method to evaluate whether our $\Omega_{\text{Si}^{\text{i+3}}}$ indicated a significant increase compared to the high-redshift observations. Since the inclusion of our $z \lesssim 1$ value resulted in a statistically significant rate of increase for $\Omega_{\text{Si}^{\text{i+3}}}$, we likely have detected a true increase of the $\text{Si}^{\text{i+3}}$ mass density at low redshift, though proof must await a larger $z \lesssim 1$ survey.

Our results indicate that any increase in $\Omega_{\text{Si}^{\text{i+3}}}$ at $z \lesssim 1$ is likely due to an increase in the number of high-column-density absorbers (i.e., shallower $\alpha_{N}$ compared to high redshift), since $dN(\text{Si}^{\text{i+3}})/dX$ is nearly constant from $z \approx 3 \rightarrow 0$. In general, the $\text{Si}^{\text{i+3}}$ mass density is dominated by the high-column-density absorbers, so even a small increase in their frequency will significantly change $\Omega_{\text{Si}^{\text{i+3}}}$.

4.4. $N(\text{Si}^{\text{i+3}})/N(\text{C}^{\text{i+3}})$

As mentioned previously, the ionic ratio $N(\text{Si}^{\text{i+3}})/N(\text{C}^{\text{i+3}})$ has been used to study the shape and/or evolution of the UVB. In order to construct a complete sample of systems with coverage of both doublets, we measured the upper limit for $N(\text{C}^{\text{i+3}})$ (or $N(\text{Si}^{\text{i+3}})$) when the doublet was not detected in association with the targeted $\text{Si}^{\text{iV}}$ (or $\text{C}^{\text{iV}}$, from Paper I) doublet but the spectral coverage existed. Due to how the $S/N$ changed throughout any spectrum, a $\text{Si}^{\text{iV}}$ targeted survey was not sufficient to define a complete $\text{C}^{\text{iV}}$ sample and vice versa. Ultimately, there were 12 detections and 12 lower limits for $z < 1N(\text{Si}^{\text{i+3}})/N(\text{C}^{\text{i+3}})$, with log $N(\text{Si}^{\text{i+3}}) > 11.9$ and log $N(\text{C}^{\text{i+3}}) > 13.37$.

In Figure 8, we compare our $N(\text{Si}^{\text{i+3}})/N(\text{C}^{\text{i+3}})$ sample with that from Boksenberg et al. (2003). We reproduced their Figure 16 (bottom panel) by summing the column densities of all components per system as given in their Tables 2–10. If there were components with upper limits for column densities, we set the total system column density to an upper limit if the components with upper limits were more than 30% of the total. If there were components with lower limits for column densities, we set the total column density to a lower limit. There was no case when these criteria conflicted. For the high-redshift sample, there were 39 detections and one upper limit for $N(\text{Si}^{\text{i+3}})/N(\text{C}^{\text{i+3}})$, for doublets with the observed low-redshift column density limits.

Since both high- and low-redshift samples contained at least one upper limit, we used survival analysis to enable those limits to contribute statistically. We used the Astronomy SURVival Analysis package (ASURV Rev. 1.3, last described in Lavalley et al. 1992) to compare the two $N(\text{Si}^{\text{i+3}})/N(\text{C}^{\text{i+3}})$ data sets. First, we tested whether the low-redshift $N(\text{Si}^{\text{i+3}})/N(\text{C}^{\text{i+3}})$ distribution shared the same parent population as the high-redshift ratios. From several univariate ASURV statistics,13 we conclude that the two populations are statistically similar (i.e., the null hypothesis cannot be ruled out with high confidence).

Next we measured the median ratio of the parent population with the Kaplan–Meier estimator.14 The estimated median of the combined low- and high-redshift data sets was $(N(\text{Si}^{\text{i+3}})/N(\text{C}^{\text{i+3}})) = 0.16$, and the 25th and 75th percentiles were 0.09 and 0.26, respectively.

Though the estimated means of the high- and low-redshift samples indicated that there should be no evolution of the ratio with redshift, we checked for a correlation.15 Once again, the null hypothesis (i.e., that there is no correlation) cannot be ruled out with high confidence.

4.5. Nature of Systems with $\text{Si}^{\text{iV}}$ and $\text{C}^{\text{iV}}$ Absorption

We concluded that there has been no evolution in $N(\text{Si}^{\text{i+3}})/N(\text{C}^{\text{i+3}})$ from $z = 4.5 \rightarrow 0$, based on our sample and that from Boksenberg et al. (2003). Next we explored what the lack of evolution means, and we began by disentangling the physics involved in the ratio $N(\text{Si}^{\text{i+3}})/N(\text{C}^{\text{i+3}})$:

$$
\frac{N(\text{Si}^{\text{i+3}})}{N(\text{C}^{\text{i+3}})} = \left( \frac{L_{\text{Si}}}{L_{\text{C}}} \right) \left( \frac{n_{\text{Si}}}{n_{\text{C}}} \right) \left( \frac{X_{\text{Si}^{\text{i+3}}}}{X_{\text{C}^{\text{i+3}}}} \right),
$$

13 For more information about the univariate analyses used here (the two Gehan’s, the Peto–Peto, and Peto–Prentice generalized tests), see Feigelson & Nelson (1985).

14 For a useful description of the Kaplan–Meier estimator in a context similar to that used here, see Simcoe et al. (2004).

15 For more information about the bivariate (Cox proportional hazard model, generalized Kendall’s tau, and Spearman’s rho) and linear regression analyses (EM algorithm and Buckley–James method) used here, see Isobe et al. (1986).
where \( L_X \) is the size of the cloud (enriched with X), along the line of sight, \( n_X \) is the volume density of element X, and \( \chi_X \) is the fraction of X ionized into ion X. The first term on the right-hand side is affected by the spatial distribution; the second, the metal abundances; and the third, the ionizing background.

The universal trend is for structure to collapse and become denser with age in a ΛCDM universe and for feedback processes to disperse and mix metals on varying scales. Therefore, more of the filamentary structure is enriched as the universe ages, though feedback may preferentially enrich voids instead of filaments, as a result of the density difference (Kawata & Rauch 2007). However, in our systems with both Si\textsc{iv} and C\textsc{iv} absorption, the absorption profiles trace each other quite well. There is no obvious system where \( N(\text{Si}^{\text{iv}})/N(\text{C}^{\text{iv}}) \) varies significantly between the components. Therefore, we infer \( L_{\text{Si}} \approx L_C \).

We know that the metallicity of the universe increases with age, but the relative abundance of silicon and carbon does not, necessarily, follow suit. Hence, the evolution in \( n_{\text{Si}}/n_C \) is unclear.

We explored the effect of the ionizing background with a suite of simple photoionization models, in order to develop our understanding of the last term in Equation (5). We used the spectral synthesis program CLOUDY v08, as last described by Ferland et al. (1998). We modeled the medium as a plane-parallel slab, ionized by the Haardt & Madau (1996, updated in 2005) UVB, for quasar (Q), galaxy (G), and quasar+galaxy (Q+G) models. We set the number density of hydrogen \( n_H = 0.1 \text{ cm}^{-3} \), though our models were insensitive to this parameter in the optically thin regime. We assumed a neutral column density \( \log N_H = 16 \) and metallicity \( Z = 0.001 Z_\odot \). We tested two cases: solar relative abundances and Si-enhanced; the increase in silicon was such that [Si/C] = +0.77, as measured in Aguирre et al. (2004). We varied the ionization parameter \( \log U = [-5, -4, -3, -2, -1] \), which is a dimensionless ratio of the flux of hydrogen-ionizing photons to \( H_\odot \cdot c \).

All models could reproduce the observed lack of evolution in the ratio of hydrogen with redshift. In Figure 9, we show \( N(\text{Si}^{\text{iv}})/N(\text{C}^{\text{iv}}) \) as a function of redshift for the three UVB models, with solar relative abundances and Si-enhanced, and for \( \log U = -2 \). We did not fit the observations, but they are reproduced well by the solar relative abundance, Q+G model. More importantly, all UVB models could reproduce a shape consistent with the observed lack of evolution in \( N(\text{Si}^{\text{iv}})/N(\text{C}^{\text{iv}}) \), given the right choices of, e.g., \( \log U \), \( \log N_{\text{H}i} \), [Si/C]. The overall magnitude is nearly freely scalable by adjusting these parameters, because the observed sample of Si\textsc{iv} and C\textsc{iv} doublets is not drawn from a single type of cloud, as we modeled. Evidently, there is no need for a particular soft UVB (i.e., model G) to reproduce the lack of redshift evolution in \( N(\text{Si}^{\text{iv}})/N(\text{C}^{\text{iv}}) \).

For there to be no evolution in \( N(\text{Si}^{\text{iv}})/N(\text{C}^{\text{iv}}) \) from \( z = 4.5 \rightarrow 0 \), the abundance, ionizing background, and structure of silicon- and carbon-enriched gas are constrained to be in “balance.” The observations indicate that these three processes evolved to maintain a constant ratio of \( (N(\text{Si}^{\text{iv}})/N(\text{C}^{\text{iv}})) \approx 0.16 \) for nearly 12 Gyr, for absorbers with \( \log N(\text{Si}^{\text{iv}}) > 11.9 \) and \( \log N(\text{C}^{\text{iv}}) > 13.37 \).

Disentangling the effect of the detailed physics (e.g., changing silicon and carbon abundances, variation in the physical properties of absorbing clouds) requires cosmo-
logical hydrodynamic simulations that could resolve enrichment processes in galactic halos and the large-scale structure.

5. SUMMARY

We conducted a blind survey for $z \lesssim 1$ Si iv doublets in the HST UV spectra of 49 quasars. We identified 22 definite Si iv systems ($G = 1$) and six “highly likely” ones ($G = 2$), and this represents the largest sample of low-redshift Si iv doublets prior to Servicing Mission 4. From a sample of 20 $z \lesssim 1$ Si iv doublets with both lines detected at $W_r > 3\sigma_W$, we measured a line density $dN_{\text{Si}^\text{IV}}/dX = 1.4^{+0.4}_{-0.3} \text{log } N(\text{Si}^+) > 12.9$.

We constructed frequency distributions of the column densities $f(N(\text{Si}^+))$ and the rest equivalent widths $f(W_r)$ prior to Servicing Mission 4. From a sample of 20 $z \lesssim 1$ Si iv doublets with both lines detected at $W_r > 3\sigma_W$, we measured a line density $dN_{\text{Si}^\text{IV}}/dX = 1.4^{+0.4}_{-0.3} \text{log } N(\text{Si}^+) > 12.9$.

In general, more observations—at low and high redshift—are needed to increase the statistical significance of the trends that are currently highly suggestive. We eagerly anticipate new low-redshift results from the HST Cosmic Origins Spectrograph (Morse et al. 1998). Meanwhile, cosmological hydrodynamic simulations should be leveraged to understand how metal production and dispersal and the ionizing background interact to evolve Si iv and C iv absorbers in tandem.

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