Integrated spectral energy distributions of binary star composite stellar populations

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ABSTRACT

This paper presents the theoretically integrated spectral energy distributions (SEDs) of binary star composite stellar populations (bsCSPs) in early-type galaxies and how the bsCSP model can be used for the spectral studies of galaxies. All bsCSPs are built based on three adjustable inputs (metallicity, ages of old and young components). The effects of binary interactions and stellar population mixture are taken into account. The results show some ultraviolet (UV) upturn SEDs naturally for bsCSPs. The SEDs of bsCSPs are affected obviously by all of the three stellar population parameters, and the effects of three parameters are degenerate. This suggests that the effects of metallicity, and the ages of the old (major in stellar mass) and young (minor) components of stellar populations should be taken into account in the SED studies of early-type galaxies. The sensitivities of SEDs at different wavelengths to the inputs of a stellar population model are also investigated. It is shown that UV SEDs are sensitive to all of the three stellar population parameters, rather than to only stellar age. Special wavelength ranges according to some SED features that are relatively sensitive to the stellar metallicity, young-component age and old-component age of bsCSPs are found by this work. For example, the shapes of SEDs with the wavelength ranges of 5110–5250, 5250–5310, 5310–5350, 5830–5970 and 20 950–23 550 Å are relatively sensitive to the stellar metallicity of bsCSPs. The shapes of SEDs within 965–985, 1005–1055 and 1205–1245 Å are sensitive to the old-component age, while SED features within the wavelength ranges of 2185–2245, 2455–2505, 2505–2555, 2775–2825 and 2825–2875 Å are sensitive to the young-component age. The results suggest that some line indices within these special wavelength ranges are possibly better for stellar population studies compared to the others, and greater weights may be given to these special SED parts in the determination of the stellar population parameters of early-type galaxies from fitting SEDs via bsCSPs.

Key words: techniques: spectroscopic – galaxies: fundamental parameters – galaxies: stellar content.

1 INTRODUCTION

The integrated spectral energy distribution (SED) is one of the most important tools for studying early-type galaxies. It can be used to investigate a lot of physical information about galaxies, because SEDs at different wavelengths are usually dominated by different physical parameters. For example, SED fitting can give an insight into the stellar mass, age, metallicity and redshift of galaxies. On the studies of SEDs of galaxies, one can use both observational spectral templates or theoretical SEDs. When theoretical SEDs are taken for studying the stellar populations of early-type galaxies, many evolutionary stellar population synthesis models, e.g. Leitherer et al. (1999) and Vázquez et al. (Starburst99; 2007), Vazdekis (1999), Schulz et al. (2002), Cerviño, Mas-Hesse & Kunth (2002), Robert et al. (2003), Bruzual & Charlot (2003), Le Borgne et al. (2004, PEGASE-HR), Maraston (2005), Lançon et al. (2008), Mollá, García-Vargas & Bressan (2009), are usually used. However, all of these models are single star simple stellar population (ssSSP) models, which do not take the effects of binary evolution and...
population mixture into account. Meanwhile, observations in the Galaxy show that more than a half of stars are in binaries. Studies on the evolution of binary stars show that they have significantly different evolution processes compared to single stars. When binary stars are used to model the SEDs of the stellar populations of early-type galaxies, some recent studies show that binary evolution can affect a number of spectral population synthesis studies significantly (van Bever et al. 1999; Van Bever & Vanbeveren 2003; Li & Han 2008a,b, 2009; Han et al. 2010; Li et al. 2010; Zhang et al. 2010), especially in the UV (e.g. Belkus et al. 2003), X-ray (Van Bever & Vanbeveren 2000) and optical bands (Han, Podsidiłowski & LymasGRAY 2007). Therefore, it is necessary to model the SEDs of galaxies via binary stars. In addition, a lot of works (e.g. Mármol-Queraltó et al. 2009) showed that galaxies, even early-type ones, should undergo more than one starburst. It means that there are multiple stellar populations in early-type galaxies. This calls for composite stellar population (CSP) models instead of simple stellar population (SSP) models to give more detailed SED studies to early-type galaxies. Han et al. (2007) tried to explain the UV upturn of SEDs of elliptical galaxies via binary stars, but all populations are assumed the solar metallicity (of SEDs of elliptical galaxies via binary stars, but all populations are assumed the solar metallicity (Z) = Z⊙). Because many works show wide metallicity ranges for early-type galaxies, it is necessary to study the SEDs of early-type galaxies using a more advanced binary star CSP (bsCSP) model, which covers a large metallicity range. This paper presents a new model, via assuming an old (major in stellar mass) and a young (minor in stellar mass) SSP component for a bsCSP. The assumption of two populations for a bsCSP is taken because early-type galaxies have relatively simple star formation histories compared to other galaxies, and it is possible to give the main information of stellar populations of early-type galaxies via two stellar population components. Our work aims to investigate how model inputs affect the SEDs of populations and to find the sensitivities of different parts of SEDs to model inputs.

This paper is organized as follows. In Section 2, we introduce the bsCSP model. In Section 3, we study how different inputs of a bsCSP model affect the SEDs of populations. In Section 4, we study the sensitivities of SEDs at different wavelengths to three main parameters of a bsCSP. In Section 5, the bsCSP model is used to determine some important parameters of three elliptical galaxies. Finally, we give our conclusion and discussion in Section 6.

2 Binary Star Composite Stellar Population Model

There are two steps to build the SEDs of bsCSPs. First, a large data base of SEDs of bsSSPs is built. Stars in a bsSSP are assumed to form at a starburst and have the same metallicity. Secondly, the SEDs of bsCSPs are built on the basis of the SED data base of bsSSPs. Because there are no common results for the stellar mixture of galaxies, we take a simple way to model bsSSPs. Each bsCSP is assumed to contain a pair of old and young components with the same metallicity. This assumption is in agreement of previous studies that early-type galaxies are dominated by old populations and there is only a little fraction of young population in such galaxies. The mass fraction of young component is assumed to be dependent on the ages of two components of bsCSPs, which is calculated by formula (1). This means that the mass fraction of the young component declines exponentially with increasing difference between the ages of old and young components. This is in agreement with some previous studies on the star formation histories of early-type galaxies, e.g. Thomas et al. (2005):

\[ F_2 = 0.5 \exp \left( \frac{t_2 - t_1}{\tau} \right), \]  

(1)

where \( F_2 \) is the mass fraction of young component; \( t_1 \) and \( t_2 \) are the ages of the old and young components of a bsCSP. As a standard model, \( \tau \) is taken as 3.02, according to the observational fraction of bright early-type galaxies with recent (≤1 Gyr) star formation at a level more than 1–2 per cent (Yi et al. 2005; Li & Han 2007).

The important ingredients and specialties of our model are as follows. An initial mass function (IMF) of Chabrier (2003) with lower and upper mass limits of 0.1 and 100 M⊙, respectively, is taken for bsCSPs. The evolution of binaries is calculated via the rapid stellar evolution code of Hurley, Tout & Pols (2002) (hereafter Hurley code), so that most of binary evolution processes such as mass transfer, mass accretion, common envelope (CE) evolution, collisions, supernova kicks and angular momentum loss are included. Different mass transfer mechanisms, i.e. dynamical mass transfer, nuclear mass transfer and thermal mass transfer are taken into account using the results of many works (e.g. Hjellming & Webbink 1987; Tout et al. 1997). One can see Hurley et al. (2002) for more details. According to a previous work, in all binary interactions, Roche lobe overflow and CE may affect the formation of hot subdwarfs and then affect the UV spectra of stellar populations more significantly (Han et al. 2007).

The default values in Hurley code, i.e. 0.5, 1.5, 1.0, 0.0, 0.001, 3.0, 190.0, 0.5 and 0.5, are taken for wind velocity factor (βw), Bondi-Hoyle wind accretion factor (αw), wind accretion efficiency factor (η), fraction of accreted material retained in supernova eruption (ϵ), CE efficiency (αCE), dispersion in the Maxwellian distribution for the supernova kick speed (σk), Reimers coefficient for mass-loss (η) and binding energy factor (κ), respectively. Although the above default values remain in somewhat large uncertainties, the results for spectral stellar population synthesis will not be affected too much by the uncertainties of these parameters. We test the effects of uncertainties in these parameters by comparing four modes with different input values. For convenience, the mode of stellar evolution by taking the default values will be called as default mode hereafter. Then the evolutionary results of other three modes (A, B and C) are compared to that of the default mode. In mode A, we take the lowest values for all parameters (i.e. 0.125, 0.0, 0.0, 0.0, −1.0, 0.5, 0.0, 0.0 for βw, αw, μw, Bw, ϵ, αCE, σk and η, respectively) to evolve 5000 binaries with metallicities (Z) of 0.01, 0.02 and 0.03 from zero-age main sequence to 15 Gyr, with an age interval of 0.1 Gyr. Then we compared the effective temperature (Teff) and surface gravity (log g) of stars in mode A to those obtained in the default mode. It shows that the two main stellar parameters (Teff and log g) change by only 0.20 per cent on average when taking mode A instead of the default mode. Then in mode B, the same star sample as in mode A is evolved, but the largest or large (for σk) values are taken for all parameters (i.e. 7.0, 2.0, 1.0, 1.0, 1.0, 10.0, 5.0, 0.0 for βw, αw, μw, Bw, ϵ, αCE, σk and η, respectively) when evolving binaries. It shows that the average change of stellar parameters is 0.48 per cent compared to the default mode, which is larger than in mode A. In the third mode, i.e. mode C, we set some random values to evolve 5000 binaries. The values for βw, αw, μw, Bw, ϵ, αCE, σk and η are set as 3.0, 1.0, 0.3, 100.0, 0.5, 5.0, 300.0 and 0.7, respectively. This mode shows an average change in stellar parameters of 0.35 per cent compared to the default mode. Because the fitted formulae used by Hurley code to evolve stars can lead to uncertainties up to about 5 per cent (Hurley et al. 2002), the changes in stellar evolutionary results that
are caused by the uncertainties in the default parameters are much less than the uncertainties that result from using fitted formulae to evolve stars. Therefore, we conclude that the uncertainties in the default values of Hurley code will not change the main results of statistical studies like spectral stellar population synthesis. Thus, we take the default values of Hurley code to evolve all stars in this work. Note that 3.0 is taken for $\alpha_{\text{CE}}$ here also because it can produce enough double degenerates, while a widely used value of 1.0 cannot when Hurley et al. (2002) tried to check the effects of the input parameters of the Hurley code via the population synthesis of birth rates and the Galactic numbers of the various types of binary.

When generating the star sample of populations, a uniform distribution is taken for the ratio $(q, 0–1)$ of the mass of the secondary to that of the primary (Mazeh et al. 1992; Goldberg & Mazeh 1994), and the mass of the secondary is then calculated from $q$ and the primary mass, where the mass of the primary is generated via a Monte Carlo technique according to the selected IMF. The separation $(a)$ of two components of a binary is generated following the assumption that the fraction of binary in an interval of log $a$ is constant when $a$ is big $(10 < a < 5.75 \times 10^6 \, R_\odot)$ and it falls off smoothly when $a$ is small $(\leq 10 \, R_\odot)$ (Han, Podsiadlowski & Eggleton 1995), which can be written as

$$ap(a) = \begin{cases} a_{\text{sep}}(a/a_0)^\psi, & a \leq a_0, \\ a_{\text{sep}}, & a_0 < a < a_1, \end{cases}$$

where $a_{\text{sep}} \approx 0.070$, $a_0 = 10 \, R_\odot$, $a_1 = 5.75 \times 10^6 \, R_\odot$ and $\psi \approx 1.2$. A uniform distribution is taken for the eccentricity ($e$) of each binary system. Most binary interactions, e.g. mass transfer, CE evolution, collisions and tidal interactions, are taken into account in the bsCSP model. The fraction of binary stars in each population (4000000 single stars) is taken as the typical value of the Galaxy, i.e. 50 per cent. When calculating the SEDs of stellar populations, the BaSeL 3.1 spectral library (Lejeune, Cuisinier & Buser 1997, 1998; Westera et al. 2002) is used because of its wide wavelength coverage and reliability. On average, the spectral library leads to about 3 per cent uncertainties in the final SEDs of bsCSPs.

3 EFFECTS OF MODEL INPUTS ON SEDs

3.1 Effect of stellar metallicity

Most previous works took the solar metallicity for studying the SEDs of early-type galaxies, especially elliptical galaxies, but the stellar metallicities ($Z$) of early-type galaxies distribute in a large range (e.g. 0.005–0.05, Gallazzi et al. 2005). Therefore, it is interesting to investigate how stellar metallicity changes the SEDs of the stellar populations of early-type galaxies. The effects of metallicity on the SEDs of bsCSPs are investigated in this section. Figs 1 and 2 show some results. In each panel of the two figures, the SEDs of bsCSPs are compared. The symbols ‘$t_1$’ and ‘$t_2$’ denote the ages of old and young components, respectively. As can be seen, stellar metallicity affects the SEDs of bsCSPs significantly. In particular, UV SED is shown to be very sensitive to stellar metallicity. When stellar metallicity changes from 0.004 to 0.03, the logarithmic flux in the UV band can change as large as 1.8 dex. Even when stellar metallicity changes only from 0.2 to 0.01, the UV spectra of bsCSPs can change a lot for most of stellar populations. This suggests that stellar metallicity should be taken into account when explaining the UV spectra of early-type galaxies.

3.2 Effect of old-component age

This section investigates how the age of old component, $t_1$, affects the SED of bsCSPs. As examples, some stellar populations with
Figure 2. Similar to Fig. 1, but for bsCSPs with a young-component age of 3 Gyr.

Figure 3. Comparison of the SEDs of bsCSPs with different old-component ages. The ages of the old components of bsCSPs, $t_2$, are set to be different (1, 2, 3 and 5 Gyr) in four panels, while the stellar metallicity, $Z$, is set to be the same. Stellar populations with different old-component ages are shown in various colours or line types.

3.3 Effect of young-component age

The effects of the age of the young component, $t_1$, on the SEDs of bsCSPs are investigated in this section. As examples, three stellar metallicities (0.01, 0.02 and 0.03) and four ages (5, 8, 10 and 13 Gyr) are taken for the stellar metallicity and old-component age of bsCSPs, respectively. The results are shown in Figs 6–8. It is...
found that the age of young component has no obvious effect on SEDs, when the young components of bsCSPs are older than 2 Gyr. However, $t_2$ affects the SEDs of bsCSPs significantly when it is less than 2 Gyr. In detail, the change in the logarithmic flux of the UV SED of bsCSP can be as large as 2 dex, when $t_2$ changes in a wide range, e.g. 0.5–10 Gyr. In addition, it is found that the younger the young component, the larger the flux in the UV band, for bsCSPs with the same metallicity ($Z$) and old-population age ($t_1$). In addition, the three figures show that the shape of the SED does not change obviously for a wavelength shorter than about 1000 Å, when $t_2$ changes from $t_1$ to 0.5 Gyr. Meanwhile, the shape of the SED from 1000 to 3000 Å is affected by $t_2$ more obviously. This confirms that the SED within a wavelength range of 1000–3000 Å is sensitive to the age of the young component. This agrees to many previous works, e.g. Yi et al. (1999) and Han et al. (2007).
Figure 6. Comparison of the SEDs of bsCSPs with the same metallicity and old-component age, but various young-component ages. In each panel, bsCSPs with various young-component ages are shown in different colours. Note that colours have different meanings in four panels.

4 SENSITIVITIES OF SEDs TO STELLAR POPULATION MODEL INPUTS

Because three input stellar population parameters (metallicity, old- and young-component ages) have some effects on the SEDs of bsCSPs, it is difficult to find SED features for determining each parameter from the results shown in Section 3. We investigate the sensitivities of SEDs to the three main input parameters of bsCSPs.

The absolute sensitivity to an input stellar population parameter is defined as the average change in logarithmic flux ($\log f_\lambda$) at a wavelength, when the input parameter changes by 10 per cent (0.0026 for $Z$, 0.7 Gyr for old-component age and 0.6 Gyr for young-component age) within its possible range. The absolute sensitivities are calculated using the SEDs of bsCSPs with four metallicities (0.004, 0.01, 0.02 and 0.03) and eight old-component ages (from 5 to 12 Gyr with a 1 Gyr interval). The ages of the young components of bsCSPs take different values from its old-component age to 1 Gyr.

Figure 7. Similar to Fig. 6, but for bsCSPs with a solar metallicity ($Z = 0.02$).
Note that because this work aims to study the SEDs of early-type galaxies, the possible ranges for the metallicities and old-component ages are taken from a result of the stellar population study of nearby early-type galaxies (e.g. Gallazzi et al. 2005).

The absolute sensitivities at all wavelengths are shown in Fig. 9. It is found that the sensitivities to three stellar population parameters differ a lot at various wavelengths. As a whole, SEDs in near-infrared and infrared bands ($\lambda \geq 8000$ Å) are more sensitive to old-component age, while UV-band SED is sensitive to all of the three stellar population parameters. In other words, old-component age affects the brightness in near-infrared and infrared bands, significantly. The brightness in the UV band is affected by all of the three stellar population parameters. It is observed that the UV flux is affected by stellar metallicity significantly. The optical SEDs of bsCSPs are less sensitive to the age of young components, compared to the stellar metallicity and age of the old components of bsCSPs. This suggests that it is difficult to determine the young-component ages of bsCSPs via optical SEDs. Therefore, the results suggest using multiple-band SEDs for well determining the stellar metallicity, old-component and young-component ages of early-type galaxies via fitting SEDs.

Although Fig. 9 tells us the absolute changes in SEDs caused from the 10 per cent change of stellar metallicity, old-component age or young-component age, it is difficult to guide many of our studies. The reason is that usually we determine the stellar population parameters via fitting some relative spectral features (or shapes) of galaxies. Such special spectral features can be found by searching for the abrupt changes with the wavelength of the relative sensitivities of SEDs to three stellar population parameters. For convenience, we use relative sensitivities instead of absolute sensitivities here. The relative sensitivities of SEDs to stellar metallicity are obtained by dividing absolute sensitivities to metallicity by 1.2. The relative sensitivities of SEDs to old-component age and young-component age are obtained by dividing absolute sensitivities by 0.7. This makes the maximum of relative sensitivities of SEDs to three stellar population parameters equal to 1 for $\lambda > 500$ Å. Then, some spectral features can be found from the relative sensitivities. For example, if the sensitivity to old-component age within a special wavelength range is significantly larger than in neighbourhood, while the sensitivities to stellar metallicity and young-component age in the same wavelength range do not have similar trend, the SED feature within the given wavelength range is relatively sensitive to old-component age. It means the relative shape of SEDs within the given wavelength range can be used as an old-component age indicator. One can check the sensitivities of SEDs within 1205–1245 Å (peak at 1225 Å) in Fig. 10 for well understanding our method. Some spectral features that are sensitive to stellar metallicity and young-component age are found in the same way. Fig. 10 shows the relative sensitivities of SEDs to three stellar population parameters and the central wavelengths of some example spectral features that are sensitive to the stellar metallicity, old-component age or young-component age of bsCSPs.
young-component age are added by 0.7 and population parameters of bsCSPs. Relative sensitivities to metallicity and (e.g. 1205 Å ≤ λ ≤ 1245 Å) in the UV band are found to be more sensitive to the old-component age. It is also found from Fig. 10 that SED features within 5110–5250, 5250–5310, 5310–5350, 5830–5970 and 20950–23550 Å are relatively sensitive to the stellar metallicity of bsCSPs. The SED features within 965–985, 1005–1055 and 1205–1245 Å are sensitive to old-component age, while those within the wavelength ranges of 2185–2245, 2455–2505, 2505–2555, 2775–2825 and 2825–2875 Å are sensitive to young-component age.

It is found that although the UV SEDs of bsCSPs are related to all of the three stellar population parameters, some spectral features (e.g. 1205 Å ≤ λ ≤ 1245 Å) in the UV band are found to be more sensitive to the old-component age. This suggests that even bsCSP models instead of ssSSP models are used for studying the stellar populations of early-type galaxies; some widely used spectral line indices, e.g. Mg$b$ Fe5270, Fe5335, can be used as metallicity indicators. In addition, SED features in a wide wavelength range of the near-infrared band seem sensitive to stellar metallicity. It implies that infrared photometry has some potential for determining the stellar metallicity of galaxies. Moreover, it is shown that the age of the young components of the bsCSPs of early-type galaxies can be estimated by some SED features within narrow wavelength ranges, e.g. 2775–2825 Å (with a peak wavelength of 2800 Å), because SEDs in these wavelength ranges are relatively sensitive to the young-component age of bsCSPs.

5 APPLICATION AND COMPARISON OF bsCSP MODELS TO GALAXIES

As a test, in this section we apply the above results, i.e. the spectral data base of bsCSPs and sensitivities of different spectral parts to the input stellar population parameters, to three elliptical galaxies, i.e. NGC 1399, NGC 1404 and NGC 2865. The stellar population parameters of the three galaxies are determined via spectral fits. The observational spectra of three galaxies are taken from a data base of the UV–optical spectra of nearby quiescent and active galaxies, which was built by Storchi-Bergmann, Calzetti & Kinney (http://www.stsci.edu/ftp/catalogs/nearby_gal/sed.html), because of its convenience. Then the original spectra were degraded to fit the resolution of BaSeL 3.1 spectra. Some bsCSP models with metallicities of 0.0003, 0.001, 0.004, 0.01, 0.02 and 0.03 are used for fitting the observational spectra, and all spectra are normalized at 5510 Å. The range of the old-population age of bsCSP models is 0.1–15 Gyr. Two methods are used to get the best-fitting results. In the first method, a homogeneous weight is given for almost all spectral parts. The spectral parts that have obvious uncertainties or poor quality in the observational spectra are given lower weights in all fits. The results show that NGC 1399, NGC 1404 and NGC 2865 have redshifts of 0.004753, 0.00650 and 0.008797, respectively. Their best-fitting stellar metallicities are shown to be 0.03, 0.03 and 0.01, respectively. Meanwhile, the best-fitting ages for the old and young populations in NGC 1399, NGC 1404 and NGC 2865 are 11.8 and 10.9 Gyr, 10.3 and 1.7 Gyr, and 14.7 and 1.1 Gyr, respectively. When various weights are given to different spectral parts in the second method, the best-fitting stellar populations of NGC 1399 and NGC 1404 are shown to be different. The metallicity of NGC 1404 is shown to be 0.02, instead of 0.03. NGC 1399 is shown to include only an old (11.8 Gyr) population, while NGC 1404 is shown to include a pair of populations with ages of 13.8 and 1.3 Gyr. The detailed comparison of the observed and best-fitting bsCSP spectra of the galaxies is shown in Figs 11 and 12. Note that in the second method, four times weights are given to three spectral parts, i.e. 1205–1245, 2775–2885 and 5110–5350 Å, according to their relative sensitivities to stellar population inputs. We see that the observational spectra of three elliptical galaxies can be fitted well using those of bsCSPs, as a whole. However, it is also seen that the spectral parts around 2000 and 3000 Å are not fitted well. This actually results from the poor quality of the observational spectra around 2000 and 3000 Å, because the part around 2000 Å is a overlap region of two observational spectra, and the second-order contamination of light causes uncertainties in the part of 3000–3200 Å. In addition, only bsCSPs with six metallicities are used for the fit. This may also contribute to the differences between the observational and best-fitting spectra.
The sensitivities of SEDs to metallicity and old-component age do not change obviously with different \( \tau \) values.

The IMF of stellar populations can also affect the sensitivities of SEDs to stellar metallicity, old-component and young-component ages. However, even other IMFs are taken for building bsCSPs, the main results of this work will not change.

This work also showed some special wavelength ranges for studying the SEDs of early-type galaxies, including studying UV-band SEDs. As a standard model, we take a formula of

\[
F_\lambda = 0.5 \exp\left[\frac{(t_\tau - t_\tau)\tau}{\tau}\right] \quad \text{with} \quad \tau = 3.02
\]

for calculating the mass fraction of the young component of a bsCSP. The value of \( \tau \) affects the sensitivities of SEDs to the young-component age of bsCSPs obviously, because it changes the mass fractions of the young components of bsCSPs directly. We checked the sensitivities of SEDs to bsCSPs by taking different values (1, 2, 3, 0.2 and 5) for \( \tau \). The results showed that the bigger the value of \( \tau \), the larger the sensitivity of UV-band SEDs to young-component age. However, the sensitivities of optical and infrared-band SEDs to the young-component age are always similar.

The effects of \( \alpha \)-element enhancement were not taken into account in this work, because the Hurley code cannot evolve stars with \( \alpha \)-element enhancement. This work has been supported by the Chinese National Science Foundation (Grant Nos 10963001 and 10978010), Yunnan Science Foundation (No. 2009CD093) and Chinese Postdoctoral Science Foundation. ZL gratefully acknowledges the support of Sino-German Center (GZ858) and K. C. Wong Education Foundation, Hong Kong. JL’s work was also supported by the Light in China’s Western Region (LCWR) (No. XBXBS2011022).

REFERENCES

Belkus H., Van Bever J., Vanbeveren D., van Rensbergen W., 2003, A&A, 400, 429
Bruzual G., Charlot S., 2003, MNRAS, 344, 1000
Cerviño M., Mas-Hesse J. M., Kunth D., 2002, A&A, 392, 19
Chabrier G., 2003, ApJ, 586, L133
Conroy C., Gunn J. E., 2010, ApJ, 712, 833
Conroy C., Gunn J. E., White M., 2009, ApJ, 699, 486
Conroy C., White M., Gunn J. E., 2010, ApJ, 708, 58
Gallazzi A., Charlot S., Brinchmann J., White S. D. M., Tremonti C. A., 2005, MNRAS, 362, 41
Goldberg D., Mazeh T., 1994, A&A, 282, 801
Han Z., Podsiadlowski P., Eggleton P. P., 1995, MNRAS, 272, 800
Han Z., Podsiadlowski P., Lynam-Gray A. E., 2007, MNRAS, 380, 1098
Han Z., Chen X., Zhang F., Podsiadlowski P., 2010, in Bruzual G., Charlot S., eds, Proc. IAU Symp. 262. Evolution of Binary Stars and Its Implications for Evolutionary Population Synthesis. Cambridge Univ. Press, Cambridge, p. 44
Hjellming M. S., Webbink R. F., 1987, ApJ, 318, 794
Hurley J. R., Tout C. A., Pols O. R., 2002, MNRAS, 329, 897
Le Borgne D., Rocca-Volmerange B., Prugniel P., Lançon A., Fioc M., Soubrin C., 2004, A&A, 425, 881
Lee H., Worthey G., Dotter A., 2009, AJ, 138, 1442
Lee, H., Vasiliev, A., 2012, ApJ, 755, 39
Lee, H., Worthey G., Dotter A., 2009, AJ, 138, 1442
Leitherer C. et al., 1999, AJS, 123, 3
Lejeune T., Cuisinier F., Buser R., 1997, A&A, 125, 229
Lejeune T., Cuisinier F., Buser R., 1998, A&A, 30, 65
Li Z., Han Z., 2007, A&A, 471, 795
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Figure 12. Similar to Fig. 11, but various weights are given to different spectral parts in the spectral fit according to the different relative sensitivities of various parts.
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