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Letter to the Editor

HerMES: Far infrared properties of known AGN in the HerMES fields

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Abstract

Nuclear and starburst activity are known to often occur concomitantly. \textit{Herschel}-SPIRE provides sampling of the far-infrared (FIR) spectral energy distributions (SEDs) of type 1 and type 2 AGN, allowing for the separation between the hot dust (torus) and cold dust (starburst) emission. We study large samples of spectroscopically confirmed type 1 and type 2 AGN lying within the \textit{Herschel} Multi-tiered Extragalactic Survey (HerMES) fields observed during the science demonstration phase, aiming to understand their FIR colour distributions and constrain their starburst contributions. We find that one third of the spectroscopically confirmed AGN in the HerMES fields have 5σ detections at 250 µm, in agreement with previous (sub)mm AGN studies. Their combined \textit{Spitzer}-MIPS and \textit{Herschel}-SPIRE colours (specifically S\textsubscript{250}/S\textsubscript{70} vs S\textsubscript{70}/S\textsubscript{24}) quite clearly separate them from the non-AGN, star forming galaxy population, as their 24 μm flux is dominated by the hot torus emission. However, their SPIRE colours alone do not differ from those of non-AGN galaxies. SED fitting shows that all those AGN need a starburst component to fully account for their FIR emission. For objects at z > 2 we find a correlation between the infrared luminosity attributed to the starburst component, L\textsubscript{SB}, and the AGN accretion luminosity, L\textsubscript{acc}, with L\textsubscript{SB} \propto L\textsubscript{acc}^{0.5}. Type 2 AGN detected at 250 μm show on average higher L\textsubscript{SB} than type 1 objects but their number is still too low to establish whether this trend indicates stronger star formation activity.

Key words. galaxies: active – galaxies: Seyfert – galaxies: star formation – infrared: general – quasars: general

1. Introduction

Active galactic nuclei (AGN) and starburst activity, both among the most energetic extragalactic phenomena, have been studied separately for decades. However, it is only in the past decade and a half, with the advent of the infrared (IR) observatories (ISO, \textit{Spitzer} and now \textit{Herschel}; Pilbratt et al. 2010) that it has become clear that the two phenomena are related, and more often than not, happen concomitantly (e.g. Schweitzer et al. 2006).

The \textit{Spitzer} IRAC and MIPS cameras sampled the spectral energy distribution (SED) of extragalactic sources exactly where the peak of the AGN dust emission is expected to occur under the paradigm of an axisymmetric dust distribution (often referred to as torus) surrounding the central super-massive black hole. The IRS spectrograph provided the necessary details of the silicate feature in emission and in absorption and we now have a better understanding of the physics of hot dust around AGN (e.g. Hao et al. 2007; Levenson et al. 2007). The peak of the cold dust emission, however, a tracer of star formation, was beyond the wavelength range explored by \textit{Spitzer} for high redshift sources and it is only now, with the advent of SPIRE (Griffin et al. 2010), that the full FIR SEDs of galaxies can be built, all the way to 500 μm.

The work presented here intends to build on the previous experience gathered with \textit{Spitzer} and BLAST. We study large samples of spectroscopically confirmed type 1 and type 2 AGN in the largest \textit{Herschel} Multi-tiered Extragalactic Survey (HerMES; Oliver et al., in prep.) fields observed during the science demonstration phase (SDP). The SEDs and IR properties of many of the objects in the samples have been studied in the past (e.g. Richards et al. 2006; Hatziminaoglou et al. 2008, 2009) and their AGN properties are well constrained. The idea is to extend such study to the larger wavelengths now observed by SPIRE in an effort to also constrain the starburst component of these objects.

\textit{Herschel} is an ESA space observatory with science instruments provided by European-led Principal Investigator consortia and with important participation from NASA.

\begin{small}
\textsuperscript{1} http://hermes.sussex.ac.uk
\end{small}
Table 1. AGN sample reference, field, type, number of objects, objects with 5σ detections at 250 µm and at both 250 and 350 µm.

| Sample          | Type | Nobj | 5σ 250 µm (and 350 µm) |
|-----------------|------|------|-----------------------|
| SDSS (LS)       | 1    | 168  | 44 (26)               |
| SDSS (FLS)      | 1    | 86   | 29 (21)               |
| Papovich et al. (2006) (FLS) | 1  | 159 | 71 (42)               |
| Lacy et al. (2007) (FLS) | 2  | 20   | 5 (2)                 |
| Martínez-Sansigre et al. (2006) (FLS) | 2  | 16   | 5 (4)                 |
| Polletta et al. (2006) (LS) | 2   | 11   | 2 (1)                 |
| Mainieri et al. (2002) (LS) | 2   | 9    | 0 (0)                 |
| **Total**       |      | 469  | 156 (96)              |

2. AGN samples, their SPIRE detections and mid-to-FIR colours

The AGN master sample considered here consists of a total of 469 spectroscopically confirmed type 1 and type 2 AGN in the Lockman-SWIRE (LS) and the Spitzer First Look Survey (FLS) fields, with redshifts that extend to z > 4. More specifically, it includes SDSS quasars in LS and FLS; mid-infrared (MIR) selected AGN in FLS with spectroscopic redshifts from the MMT-Hectospec (Papovich et al. 2006); two MIR-selected type 2 AGN samples from Martínez-Sansigre et al. (2006) and Lacy et al. (2007); and two X-ray selected type 2 samples from Mainieri et al. (2002) and Polletta et al. (2006). Their SPIRE fluxes are estimated via linear inversion methods, using, whenever available, the positions of known 24 µm sources as priors (see Roseboom et al., in prep.). The details of the various subsamples and the numbers of 250 µm and additional 350 µm 5σ detections are given in Table 1. Note that none of the X-ray selected type 2 AGN from Mainieri et al. (2002) were detected by SPIRE. All 5σ detections at 250 µm and 350 µm have fluxes above 12.8 mJy and 12.2 mJy, respectively. No flux cut is applied at 500 µm, these fluxes are used, instead, at face value. The different detection rates of the various samples in the two fields are the result of the different depths (FLS being deeper than LS, Oliver et al. 2010, this volume) as well as the selections of the various samples.

It is worth noting that one third of the objects have 5σ detections at 250 µm, also reported by Elbaz et al. (2010). This detection rate is in excellent agreement with results obtained from the studies of bright (MB < −26.1), high redshift (z ≥ 1.8) quasars at 1.2 mm, with the MAX-Planck Millimeter Bolometer (MAMBO) array at IRAM 30-m telescope (Carilli et al. 2001; Omont et al. 2001). It is, however, higher than the ~5%, reported by Priddey et al. (2003) for 1.5 < z < 3.0 radio-quiet, luminous (MB < −27.5) quasars, observed at 850 µm with SCUBA on the James Clerk Maxwell Telescope (JCMT). Even though our derived detection rate holds for the whole absolute magnitude range of our sample (~19 > MB > −28.3) as well as in broad MB bins (here we use MB, as the SDSS g-band filter is the closest to the B-band used in the other studies), if we look at the more rare, brightest objects (MB < −27.5), we only find about 20% of them having 5σ detections at 250 µm, close to the 15% reported by Priddey et al. (2003).

Though it is difficult to say anything about the significances of the detections of the rest of the objects in an objective way, the observed properties of the AGN with and without SPIRE counterparts are quite similar, both in terms of optical and IR fluxes and in redshift. Therefore, since the present study only makes use of objects with 5σ detections at 250 µm, and since 250 µm emission arises from cold dust in the star forming regions, the 250 µm-detected sample may be biased towards objects with stronger star formation activity. The issue of SPIRE detection rates of AGN is addressed in more detail in Stevens et al., in prep.

Figure 1 shows the MIPS-SPIRE colours S250/(S70 + S250) (a) and SPIRE colours S350/(S250 + S350) vs. S350/S250 (b) of all AGN with 250 µm and 350 µm detections above 5σ (type 1: circles; type 2: squares) colour-coded based on their redshift, and compared to all SPIRE sources in the FLS field at the same SPIRE detection limits (black dots). AGN have much bluer colours and tend to separate nicely from the bulk of star forming non-active galaxies, as their 24 µm fluxes are dominated by the hot dust (torus) emission. This is in agreement with recent MIR spectroscopic studies with Spitzer/IRAC showing that the MIR quasar emission mostly arises from the torus, independent on their FIR properties (Netzer et al. 2007). The presence of strong PAH features, however, may slightly weaken this effect in certain redshift ranges. The 7.7 µm PAH feature, in particular, might affect the 24 µm flux of z ~ 2 AGN, when present (Lutz et al. 2008).

In the SPIRE bands the colours of AGN are indistinguishable from those of the star forming non-active galaxies, suggesting that AGN actually appear like starburst galaxies in the FIR. This is consistent with longer wavelength studies that show the FIR emission of submm-luminous quasars and their starburst to share many properties of submm galaxies. Their FIR luminosities are similar, and the best evidence is probably that they follow practically the same FIR/CO luminosity relation as submm galaxies and local ULIRGs (see e.g. Fig. 5 of Riechers et al. 2006).
A variety of multi-wavelength datasets covering the energy range from the radio to the X-rays is available for the objects in our master sample, but in order to have a uniform wavelength SED coverage to perform the fitting, we only used the available photometry from SDSS DR7 (Abazajian et al. 2009), 2MASS and 2MASSx6 in the Lockman field (Beichman et al. 2003), Spitzer IRAC and MIPS data from the SWIRE (Lonsdale et al. 2004) and FLS Spitzer surveys, and new SPIRE HerMES data. PACS (Poglitsch et al. 2010) data have not been used, as they are only available for parts of the HerMES SDP fields.

3. Star formation in AGN

The observed SEDs, described in Sect. 2, were compared, by means of SED fitting with a standard $\chi^2$ minimisation, to a series of models comprising three different components: a stellar component composed by various simple stellar population (SSP) models build using the Padova evolutionary tracks (Bertelli et al. 1994); a grid of AGN/torus models that include both a toroidal and a flared disk dust geometry presented in Fritz et al. (2006); and six empirical starburst SEDs. For a full description of the SED fitting and individual model components see Hatziminaoglou et al. (2008, 2009). For reasons explained in detail in Hatziminaoglou et al. (2008) the reduced $\chi^2$ can reach high values without undermining our confidence in the fits. We will restrict the present study to the objects with fits having reduced $\chi^2 < 10$.

This leaves a total of 68 (42) type 1 and 11 (7) type 2 AGN, with 5$\sigma$ detections at 250 $\mu$m (and 350 $\mu$m), respectively. As already mentioned, the 500 $\mu$m are taken at face value, even though the detection level of about one third of them falls below 2$\sigma$. The number drops to about 15% for the objects with a 5$\sigma$ detection at 350 $\mu$m. Despite the low significance of some of the 500 $\mu$m data points, they follow nicely the observed SEDs as traced by the other FIR points and are unlikely to affect the fit, because of their large photometric errors. SPIRE 500 $\mu$m nondetections are not treated as upper limits and are excluded from the fits. Examples fits for a type 1 and a type 2 AGN are shown in Fig. 2.

As a first remark we point out that, in order to reproduce the observed SPIRE data points, a starburst template is always needed, even if we allow for very large (kpc-scale) tori. Large tori are, in any case, unphysical in the sense that they extend well into the host galaxy where other physical phenomena such as star formation may occur, and the AGN is no longer the primary source of dust heating. The SED fitting results in the set of estimations of physical parameters, describing the various components. Here we will focus on the accretion luminosity, $L_{\text{acc}}$, the model luminosity of the accretion disk ranging from soft X-rays to the optical wavelengths scaled to the observed data points, and the IR luminosity of the starburst component, $L_{\text{SB}}$, integrated between restframe 8 and 1000 $\mu$m, as a direct measure of star formation.

Figure 3 shows $L_{\text{SB}}$ as a function of $L_{\text{acc}}$ with each point colour-coded according to redshift for 250 $\mu$m 5$\sigma$ detections alone (open symbols) and additional 350 $\mu$m detections (filled symbols). A broad correlation of $L_{\text{SB}}$ can be seen with $L_{\text{acc}}$, both quantities, however, also scale with redshift, as seen from the colour-coding of the points.

If we divide the sample in bins of redshift, the picture becomes less clear. Table 2 shows the Pearson correlation coefficients for the quantities $L_{\text{SB}}$ and $L_{\text{acc}}$ in five redshift bins. (The missing value in the last column reflects the very small number of points in the relevant bin, for which a Pearson correlation would be meaningless.) The very low values of the coefficients suggest that the observed global trend may reflect the fact that the sample being flux limited, more distant objects will also be intrinsically more luminous. A lack of correlation between $L_{\text{SB}}$ and $L_{\text{acc}}$ would imply that star formation activity is not influenced by the presence of an active nucleus in the centre of the galaxy. Nevertheless, since the derived $L_{\text{acc}}$ only covers at most two orders of magnitude at any given redshift, and since we are dealing with a limited number of objects, it may be that our sample is not adequate to detect any such correlations. Considering the $z > 2.0$ bin alone we do find a 95% (90%) probability of $L_{\text{SB}}$ and $L_{\text{acc}}$ correlating as $L_{\text{SB}} \propto L_{\text{acc}}^{0.35}$, for objects with 5$\sigma$ detections at 250 $\mu$m (and additionally at 350 $\mu$m). This value is in surprisingly good agreement with that found by Wang et al. (2008) in their study of $z \sim 6$ quasars with MAMBO.

An important issue in the study of nuclear and star formation activities is that of the occurrence of star formation in the...
Table 2. Correlation coefficients $r$ and $r'$ between $L_{SB}$ and $L_{acc}$ in bins of objects considering the subsamples with 5$\sigma$ detections at 250 $\mu$m alone (N) and those with both 250 and 350 $\mu$m 5$\sigma$ detections (N'), respectively.

| $z$ | N  | N' | r    | r'   |
|-----|----|----|------|------|
| < 0.5 | 15 | 4  | 0.063 |      |
| 0.5 < $z$ < 1.0 | 14 | 10 | 0.073 | 0.057 |
| 1.0 < $z$ < 1.5 | 9  | 6  | 0.054 | 0.241 |
| 1.5 < $z$ < 2.0 | 21 | 13 | 0.109 | 0.241 |
| $z$ > 2.0 | 18 | 16 | 0.496 | 0.468 |

Fig. 4. Distribution of $L_{SB}$ (integrated between 8 and 1000 $\mu$m), as an indicator of star formation, for type 1 (black) and type 2 (grey) AGN with 5$\sigma$ detections at 250 $\mu$m.

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