Structure and kinematics of Type II Cepheids in the Galactic bulge based on near-infrared VVV data

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ABSTRACT

Context. Type II Cepheids (T2Cs) are radially pulsating variables that trace old stellar populations and provide distance estimates through their period-luminosity (PL) relation.

Aims. We trace the structure of old stellar population in the Galactic bulge using new distance estimates and kinematic properties of T2Cs.

Methods. We present new near-infrared photometry of T2Cs in the bulge from the VISTA Variables in the Vía Láctea survey (VVV). We provide the largest sample (894 stars) of T2Cs with JHK observations that have accurate periods from the Optical Gravitational Lensing Experiment (OGLE) catalog. Our analysis makes use of the K-band time-series observations to estimate mean magnitudes and individual distances by means of the PL relation. To constrain the kinematic properties of our targets, we complement our analysis with proper motions based on both the VVV and Gaia Data Release 2.

Results. We derive an empirical K-band PL relation that depends on Galactic longitude and latitude: $K_{0} = (10.66 \pm 0.02) - (2.21 \pm 0.03) \times \log P + (-0.020 \pm 0.003) \times l + (0.050 \pm 0.008) \times |b|$ mag; individual extinction corrections are based on a 3D reddening map. Our targets display a centrally concentrated distribution, with solid evidence of ellipsoidal symmetry – similar to the RR Lyrae ellipsoid and a few halo outliers up to 100 kpc. We obtain a distance from the Galactic center of $R_{0} = 8.46 \pm 0.03 \text{(stat.)} \pm 0.11 \text{(syst.)} \text{kpc}$. We also find evidence that the bulge T2Cs belong to a kinematically hot population, as the tangential velocity components $(\sigma_{\phi}, \sigma_{h}) = (104.2 \pm 3.0 \text{ km s}^{-1}, 96.8 \pm 5.5 \text{ km s}^{-1})$ agree within 1.2σ. Moreover, the difference between absolute and relative proper motion is in good agreement with the proper motion of Sgr A* from VLBA measures.

Conclusions. We conclude that bulge T2Cs display an ellipsoidal spatial distribution and have kinematics similar to RR Lyrae stars, which are other tracers of the old, low-mass stellar population. T2Cs also provide an estimate of $R_{0}$ that agrees excellently well with the literature, taking account of the reddening law.

Key words. stars: variables: Cepheids – Galaxy: bulge – Galaxy: structure – Galaxy: kinematics and dynamics

1. Introduction

In the Galactic bulge, the Red Clump (RC) stars, which are core-helium burning low-mass stars with ages from intermediate ($1 \leq \text{age} \leq 10 \text{ Gyr}$) to old ($\geq 10 \text{ Gyr}$) and average to high metallicities ([Fe/H] $\geq -1.5$ dex, Cole 1998; Hill et al. 2011), are used extensively to study their kinematics, chemical abundances, and spatial distributions (McWilliam & Zoccali 2010; Saito et al. 2011; Gonzalez et al. 2015; Zoccali et al. 2017). High-resolution spectroscopic studies of RCs suggest that only stars that are more metal-rich than [Fe/H] $\sim -0.5$ dex trace an X-shaped structure that appears to be in the form of a peanut/bulge bulge (Ness et al. 2013; Zoccali & Valenti 2016). Recently, mid-infrared data obtained with the Wide-field Infrared Survey Explorer (WISE) also showed a clear large-scale X structure (Ness & Lang 2016). In contrast, stars that are more metal-poor than [Fe/H] $\sim -0.5$ dex not only display a centrally concentrated axisymmetric spatial distribution, but also reveal different kinematics (Ness et al. 2013; Valenti et al. 2016; Zoccali et al. 2017).

However, the spatial distribution of the old and more metal-poor population of stars in the bulge, which is traced by RR Lyrae (RRLs), is still under discussion. Pietrukowicz et al. (2015) found an ellipsoidal distribution (axis ratios of 1:0.49 ± 0.03) or via anonymous ftp to cdsarc.u-strasbg.fr or via http://cdsarc.u-strasbg.fr/viz-bin/qcat?J/A+A/619/A51

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2. Data

Light curves. We used the aperture photometry of VVV DR4 data (Minniti et al. 2010; Saito et al. 2012) that is publicly available through VISTA Science Archive (VSA)\(^1\). We compared point spread function (PSF) with aperture photometry for a sample of our targets and found the differences to be negligible, but aperture photometry has the advantage of being available in the entire VVV survey area. As a first step, we matched the OGLE IV catalog of T2Cs (Soszyński et al. 2017) with the source detection catalog of VVV, adopting a matching radius of 2\(\prime\). This allowed us to retrieve 894 of 924 targets within the VVV survey area. A posteriori, we checked that all the good matches are within 1\(\prime\) of the OGLE coordinates. Of those that were not retrieved, 25 are outside the VVV area, and for five of them we could not find a good match, even with a larger searching radius of 10\(\prime\). Of these 894 T2Cs, according to the classification of Soszyński et al. (2017), 369 are BL Herculis (BLHs), 343 are W Virginis (WVs), 28 are peculiar W Virginis (pWVs), and 154 are RV Tauri (RVTs). We discuss the different types of T2Cs in more detail in Sect. 4.2.

We collected both the single epoch \(JH\)-band photometry and the \(K_s\)-band time series. \(Z\) photometry was neglected because it is not useful for our goals and we cannot even estimate mean magnitudes in these bands, since light-curve templates (Bhardwaj et al. 2017a) are available only for the \(JHK_s\) bands. The number of valid \(K_s\)-band phase points per variable (those with good photometric solution) ranges from four to 185, with a median of 1. Only four variables have fewer than ten phase points, and 90% of the variables have more than 47 observations, with a good phase coverage over the whole range of periods of our targets (1–85 days). The number of available \(K_s\) epochs varies across the VVV survey area, with the majority of the VVV pointings (so-called tiles) having between 50 and 100 epochs. Given the overlap between the adjacent tiles, which amounts to about 1\(\circ\) on a side, a small fraction of our variables (34) were observed in two tiles, with up to 185 data points.

Reddening. We adopted the two Galactic bulge \(E(J-K_s)\) reddening maps of Gonzalez et al. (2012) and Schultheis et al. (2014), henceforth, G12 and S14. The reasons are manifold: (1) They were both obtained with VVV data, therefore no photometric system conversion is needed. (2) The map of G12 provides very high resolution especially in the central regions (2\(\times\)2 arcmin for \(-3.5 < b < 5\,0\), 4\(\times\)4 arcmin for \(-7.0 < b < -3.5\) and 6\(\times\)6 arcmin for \(-10.0 < b < -7.0\)), while the "pixels" of the S14 map have a size of 6\(\times\)6 arcmin everywhere. However, the map by S14 has the key advantage to be three-dimensional. It provides a grid of \(E(J-K_s)\) for 21 bins of distance, from 0 to 10.5 kpc. This is crucial for studying a complex structure such as that of the Galactic bulge and possibly intervening thick-disk populations. In the following, we indicate as \(E(J-K_s)\)\(_{G12}\) and \(E(J-K_s)\)\(_{S14}\) the two reddening values obtained from the G12 and S14 maps, respectively. As explained in Sect. 3, we adopt the 3D reddening map of S14 for our final estimates, but G12 serves as a comparison.

Proper motions. We retrieved relative proper motions for 894 of 924 targets from the publicly available VIRAC catalog (Smith et al. 2018), obtained with VVV data. According to their recommendations, we discarded all targets for which the flag unreliable, based on the validation of the photometric solutions, is equal to zero. We point out that although the cross-match was

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\(^1\) vsa.roe.ac.uk/index.html
made by unique VVV ID and not by coordinate, we found multiple (either double, triple, or quadruple) records for 120 targets in VIRAC. We checked the multiple identifications one by one on the basis of right ascension (α), declination (δ), and Ks-band magnitude. About two-thirds of the time, all of the records of a multiple identification had \textit{reliable} = 0 and were thus all discarded. Of the remaining fraction, an a posteriori check revealed that the majority had \textit{reliable} = 1 for the correct match and \textit{reliable} = 0 for the incorrect matches. This validates our selection of the best match.

We also retrieved relative proper motions for 416 targets from PSF photometry, obtained from VVV data by the method explained in \textcite{Contreras Ramos et al. 2017}. The match was performed using a searching radius of 2′′. This catalog of proper motions does not cover the entire VVV area, but only the low latitudes (−3:0 ≤ b ≤ 3:0), that is, the most crowded region, where PSF photometry has several advantages over aperture photometry. A comparison between the two sets of proper motions from these VVV is performed in Sect. 5.

Finally, we searched our targets within the recent \textit{Gaia} DR2 (\textcite{Gaia Collaboration 2016, 2018}), using a searching radius of 4′′. We retrieved matches for 920 targets. Because of the density of \textit{Gaia} DR2, a search radius of 4′′ means multiple records for almost all our targets. We note that for \textit{Gaia}, we needed a larger searching radius because α and δ are at the epoch J2015.5. We selected the best matches on the basis of the separation in (α,δ) and the V − G and I − G color indexes, and retrieved 914 targets (for the remaining 6 targets, it was not possible to select a best match based on the adopted criteria). Of these, 868 have a five-parameter \textit{Gaia} solution (coordinates, parallaxes, and absolute proper motions).

### 3. Light curves and properties of T2Cs

We have phased the VVV light curves using the periods provided by OGLE IV (\textcite{Soszyński et al. 2017}). We visually inspected all the VVV light curves and separated promising from noisy or poorly sampled light curves. We estimated the uncertainty on the mean magnitude ($eK_s$) as the sum in quadrature of the standard error of mean of the phase points around the fit (see below) plus the median photometric error of the phase points. We note that the standard error on the mean and median photometric error represents a statistical and systematic measure of uncertainties, respectively,

$$eK_s = \sqrt{\frac{1}{n} \sum_{i=1}^{n} (\text{mag}_{i} - \text{mag}_{\text{fit}})^2} + \text{median(} \text{err} \text{)}^2.$$  

We adopted $eK_s$ as a quantitative criterion to select good-quality light curves, with a threshold at $eK_s < 0.05$. We performed visual inspection to validate this threshold and further selected promising light curves within the range $0.05 < eK_s < 0.10$ mag. However, some of the brighter stars ($K > 11.5$ mag) were classified as poor-quality light curves in this range (see below), and all stars with $eK_s > 0.10$ mag were also included in the poor-quality sample. Examples of good- and poor-quality light curves are displayed in Fig. 1.

For the 161 targets with poor-quality light curves, we adopted for the mean magnitude the median of the magnitudes, converted into flux, of the individual phase points. The uncertainty on the mean magnitude is the standard deviation of the median. We do not provide a light-amplitude estimate for these targets since the uncertainty would be on the same order as the amplitude itself. We point out that most of these targets are very bright (135 of 161 have ($K_s < 11.5$ mag), meaning that they are either saturated or in the nonlinear regime of the camera. However, we also point out that not all stars in this magnitude range have poor-quality light curves. The cause of poor photometry is a synergy between saturation and crowding by other nearby bright stars. The magnitude for saturated sources is derived from the most external apertures (those with the largest radii). When another bright source contaminates the external aperture, the photometric solution is worse. In these cases, PSF photometry would not help to improve the photometry because it would be even more affected by saturation.

To compute the mean magnitudes, amplitudes, and relative uncertainties of the good-quality targets, we adopted the fits of the light curves obtained with the PLOESS \textcite{Braga et al. 2018} fitting method, which is a variant of the GLOESS method \textcite{Persson et al. 2004}. The mean magnitude is the integral of the fitting curve converted into flux, while the amplitude is the difference between the brightest and faintest points of the fit. The uncertainty on the mean magnitude was defined before as $eK_s$. The uncertainty on the amplitude was derived as the sum in quadrature of the median photometric errors of the phase points around the maximum and minimum, plus the standard deviation of these phase points around the fit of the light curve. The final value was weighted with the number of phase points around the maximum and the minimum. We also derived mean J- and H-band magnitudes by applying the light-curve templates of \textcite{Bhardwaj et al. 2017a} to the single J and H phase points and using the time of maximum from OGLE \textcite{Soszyński et al. 2017}. The photometric properties of our final sample are listed in Table 1.

Next, we corrected the mean magnitudes for extinction. As stated in Sect. 2, we derived two extinction values from two different reddening maps. We obtained $E(J − K_s)_{G12}$ and $A_{K_s(G12)}$ for each target from the online tool BEAM\textsuperscript{2}. However, estimating $E(J − K_s)_{S14}$ is not straightforward because the distance of the target needs to be known in order to obtain $E(J − K_s)_{S14}$. Distance and reddening in S14 are degenerate. Therefore, we adopted the following method. First, we located the four pixels of the map that are closest to the position of our target. We weighted the $E(J − K_s)_{S14}$ of each pixel by their inverse angular distance of the pixel center from our target. We repeated this for each of the 21 bins and obtained 21 possible values for $E(J − K_s)_{S14}$ for each target. Using an iterative method (\textcite{Bhardwaj et al. 2017b}), we simultaneously found the most plausible values of $E(J − K_s)_{S14}$ and $A_{K_s(S14)}$, and an estimate of the distance.

We adopted a new reddening law with both reddening maps to derive the extinction $A_{K_s}$ from $E(J − K_s)$. This law was derived by \textcite{Alonso-García et al. 2017} using VVV data for the innermost regions of the bulge between $|b| < 2.7$ and $|b| < 1.55$. Specifically, we adopted as the ratio of total-to-selective extinction $R_{IK} = E(J − K_s) / A_{K_s}$, the values shown in their Table 2, according to the quadrant, within their surveyed sky area. For targets outside the quoted area, we adopted $R_{IK} = 0.428$, which is the average suggested by \textcite{Alonso-García et al. 2017}.

We checked the differences $\Delta A_{K_s} = A_{K_s(G12)} − A_{K_s(S14)}$ for all our targets, and found that $\Delta A_{K_s}$ values follow an almost Gaussian distribution with a mean of $−0.007$ mag and $\sigma = 0.032$ mag. A tail of targets with $0.1 < \Delta A_{K_s} < 0.25$ mag and one target with $\Delta A_{K_s}$ as high as $0.63$ mag were found. Almost all of these targets are located at distance moduli (DM) smaller than 14.5 mag.

\footnote{\texttt{mill.astro.puc.cl/BEAM/calculator.php}}
Fig. 1. Left: example of a good-quality light curve. Right: example of a poor-quality light curve. The periods are labeled at the top of the panels, while the names, as provided in the OGLE catalog, are labeled at the bottom. The PLOESS fit of the good-quality light curve is displayed as a red line.

Table 1. Photometric properties of target T2Cs.

| ID (OGLE IV) | ID (VVV) | ID (Gaia) | Type | Period (days) | (V) mag | (I) mag | (J) mag | (H) mag | (Ks) mag | Amp(Ks) |
|--------------|----------|-----------|------|---------------|---------|---------|---------|---------|----------|--------|
| 001          | ...      | ...       | BLHer | 3.998508      | 15.759  | 14.176  | ...     | ...     | ...      | ...    |
| 002          | 515001356315 | 5990064527510861824 | BLHer | 2.284194      | 15.188  | 13.909  | 12.976 ± 0.006 | 12.644 ± 0.007 | 12.507 ± 0.004 | 0.212 ± 0.013 |
| 003          | 515001679485 | 5990903804799447004 | BLHer | 1.484493      | 16.519  | 15.061  | 14.027 ± 0.007 | 13.648 ± 0.009 | 13.409 ± 0.008 | 0.334 ± 0.019 |
| 004          | ...      | 410738597338284672 | BLHer | 1.2119999     | 16.404  | 14.856  | ...     | ...     | ...      | ...    |
| 005          | 515594023082 | 598029157674487080 | BLHer | 2.0075505     | 18.666  | 16.842  | 15.485 ± 0.024 | 14.993 ± 0.034 | 14.782 ± 0.034 | 0.407 ± 0.040 |
| 006          | ...      | 4107420780154292480 | pWVir | 7.6364832     | 14.728  | 13.352  | ...     | ...     | ...      | ...    |
| 007          | 515543342733 | 405972396740384128 | BLHer | 1.8174741     | 17.452  | 15.530  | 13.696 ± 0.010 | 13.559 ± 0.012 | 13.292 ± 0.011 | 0.286 ± 0.014 |
| 008          | 515520862858 | 405946032759214976 | BLHer | 1.1895551     | 17.765  | 15.970  | ...     | 14.034 ± 0.015 | 13.967 ± 0.017 | 0.183 ± 0.019 |
| 009          | 51555436344 | 40595044719073928 | BLHer | 1.8960480     | 17.620  | 15.657  | 14.003 ± 0.009 | 13.437 ± 0.011 | 13.266 ± 0.013 | 0.336 ± 0.029 |
| 010          | ...      | 41099654299082432 | BLHer | 1.9160655     | 16.679  | 14.969  | ...     | ...     | ...      | ...    |

Notes. Only the first 10 of the 924 lines of the table are shown. The full table is available at the CDS. (a)The full name is OGLE-BLG-T2CEP-XXX, where “XXX” is the ID appearing in the first column.

The stars are closer than the Galactic center. On the other hand, targets with ΔKs < −0.1 mag are mostly located at DMs larger than 14.5 mag. This is expected when comparing a reddening map fixed at ~8 kpc (G12) with one that takes distance into account (S14). This also suggests that we can adopt the S14 map to derive our final results.

However, since reddening and distance were derived simultaneously, AχKs may depend on the calibration selected for the PL relation. We checked that for different calibrations (see Sect. 4.2), the differences between AχKs from different calibrations are within 0.05 mag and are smaller than 0.01 mag for ~80% of the targets.

4. PL relations and distances

4.1. Empirical PL relation

Figure 2 shows the targets in the log P−K0 plane. The BLHs and WVs were dereddened adopting the S14 map, but AχKs values and distances were estimated simultaneously. Since we did not estimate the distances of pWVs and RVTs, as explained below, these targets were dereddened with AχKs, which is independent of distance.

In stellar system like the bulge, where stars are not all at the same distance (as in the case of globular clusters, external galaxies, etc.), extinction and distance are degenerate and a simple empirical PL does not provide precise insight into the structure, especially if the reddening is not constant. Moreover, there is a debate on whether the old population in the bulge is indeed spheroidal (Dékány et al. 2013; Kunder et al. 2016) or if it is ellipsoidal and tilted, similarly to the Galactic bar (Pietrukowicz et al. 2015). Therefore, as done before by Groenewegen et al. (2008), we fit the PL relation by adding the dependence on the Galactic longitude (ℓ) and latitude (b). We selected all the BLHs and WVs from our sample, and after an iterative rejection of outliers at more than 3σ, we found K0 = (10.66 ± 0.02)−(2.21 ± 0.03)·(log P − 1.2)−(0.020 ± 0.003)−(0.050 ± 0.008)·|b| mag, with a standard deviation of 0.07 mag. The positive coefficient in |b| means that fainter stars are located at higher distances from the Galactic plane, where the reddening is lower. This is an evidence that the T2C sample is biased by reddening. On the other hand, the non-zero dependence of K0 on ℓ indicates that the T2C ellipsoid is tilted. If we ignore the ℓ and |b| terms, the simple PL relation is K0 = (10.76 ± 0.02)−(2.23 ± 0.03)·(log P − 1.2) mag, and the standard deviation increases to 0.28 mag. Figure 2 shows our
newly derived empirical, coordinate-independent PL as a black solid line. We set the zero of the independent variable at log $P = 1.2$ to facilitate comparison with Groenewegen et al. (2008). The coefficients of both the coordinate-dependent and the simple PL relations agree remarkably well with those obtained with an identical approach by Groenewegen et al. (2008): all of them agree within 1σ. However, while the error on the coefficient of the $l$ coordinate in Groenewegen et al. (2008) was larger than the value itself ($-0.028 \pm 0.031$ mag$^{-1}$, thus including the zero value within 1σ), our coefficient is more precise ($-0.019 \pm 0.003$ mag$^{-1}$) because the set of variables is much larger, and it clearly rules out a PL relation that is independent of $l$. We also compared our empirical PL with that of Bhardwaj et al. (2017b), $K_0 = (10.749 \pm 0.0056) - (2.189 \pm 0.032) \cdot (\log P - 1.2)$ mag. We adjusted their zero-point since they had adopted log $P = 1.0$ as the zero of their independent variable. The slope and zero-point both agree with ours within 1σ.

The bottom panel of Fig. 2 also shows PLs from the literature, Galactic globular clusters (Matsunaga et al. 2006), Large Magellanic Cloud (LMC; Bhardwaj et al. 2017a) as purple, red, and green dashed lines, respectively. The PL relations of Matsunaga et al. (2006, 2011) and theoretical (Bono et al. 1997; Di Criscienzo et al. 2007) arguments. The advantage of adopting a metal-independent PL relation to calibrate distances is straightforward because the metallicity of our targets is unknown. Since T2Cs belong to the same stellar population as RRLs, their

| ID | $A^a_{K_{0(12)}}$ | $A^a_{K_{0(14)}}$ | $x_{\text{GAL}}$ | $y_{\text{GAL}}$ | $z_{\text{GAL}}$ | $d$ |
|----|------------------|-----------------|----------------|-------------|-------------|-----|
| 001 | ... | ... | ... | ... | ... | ... |
| 002 | $0.07 \pm 0.04$ | $0.11 \pm 0.01$ | $7.44 \pm 0.62$ | $-1.06 \pm 0.09$ | $0.62 \pm 0.05$ | $7.54 \pm 0.63$ |
| 003 | $0.12 \pm 0.04$ | $0.16 \pm 0.01$ | $9.09 \pm 0.77$ | $-1.23 \pm 0.10$ | $0.70 \pm 0.06$ | $9.20 \pm 0.78$ |
| 004 | ... | ... | ... | ... | ... | ... |
| 005 | $0.17 \pm 0.04$ | $0.21 \pm 0.01$ | $19.25 \pm 1.65$ | $-2.20 \pm 0.19$ | $1.30 \pm 0.11$ | $19.42 \pm 1.66$ |
| 006 | ... | ... | ... | ... | ... | ... |
| 007 | $0.21 \pm 0.05$ | $0.25 \pm 0.03$ | $9.15 \pm 0.78$ | $-0.43 \pm 0.04$ | $0.63 \pm 0.05$ | $9.18 \pm 0.79$ |
| 008 | $0.17 \pm 0.04$ | $0.20 \pm 0.03$ | $10.55 \pm 0.91$ | $-0.31 \pm 0.03$ | $0.83 \pm 0.07$ | $10.58 \pm 0.91$ |
| 009 | $0.20 \pm 0.05$ | $0.21 \pm 0.03$ | $9.36 \pm 0.80$ | $-0.48 \pm 0.04$ | $0.59 \pm 0.05$ | $9.39 \pm 0.80$ |
| 010 | ... | ... | ... | ... | ... | ... |

Notes. Only the first 10 of the 924 lines of the table are shown. The full table is available at the CDS. (a) For some coordinates, the G12 map does not provide an error on $E(J - K_s)$, therefore there is no error on $A^a_{K_{0(12)}}$.

4.2. Individual distances

Before deriving individual distances, we discuss a few key points below.

(1) Metallicity dependence. There is a general consensus about the independence, or a very mild dependence, less than 0.1 mag dex$^{-1}$ (Di Criscienzo et al. 2007), of the PL relation of T2Cs on metal abundance, based on both empirical (Matsunaga et al. 2006, 2011) and theoretical (Bono et al. 1997; Di Criscienzo et al. 2007) arguments. The advantage of adopting a metal-independent PL relation to calibrate distances is straightforward because the metallicity of our targets is unknown. Since T2Cs belong to the same stellar population as RRLs, their
metallicity dispersion should be similar, and it should be fairly small (0.25 dex, Pietrukowicz et al. 2012).

(2) RV Tauri and peculiar W Virginis. T2Cs are separated into BLHs (1 day $< P < 5$ days), WVs (5 days $< P < 20$ days), and RVTs ($P > 20$ days). The thresholds are those by Soszyński et al. (2011). They are based on the period distribution of OGLE III T2Cs, and the authors kept them unchanged in the OGLE IV catalog (Soszyński et al. 2017). However, there is a debate as to whether RVTs obey the same PL as shorter-period T2Cs: Matsuııga et al. (2009) using data from the InfraRed Survey Facility (IRSF) and Rippeıı et al. (2015) using VMC data, showed that in the LMC, RVTs are overluminous with respect to the extrapolation at long periods of the PL of BLHs and WVs. In contrast, Bhardwaj et al. (2017a) used NIR data in the central bar of the LMC (Macıı et al. 2015) and found that RVTs also fall on the linear PL fit to short-period T2Cs in the LMC. However, their photometry for short-period variables (that is, at fainter magnitudes) was more prone to be affected by crowding in the central regions than longer-period variables. On the other hand, RVTs in Globular clusters also follow the same PL relation as shorter period T2Cs (Matsuııga et al. 2006). We do not discuss this in detail since this is not the aim of this work, but it is crucial to remember that in addition to these empirical findings, there are two different evolutionary channels from which RVTs are generated: either from low-mass ($< 0.5 M_\odot$), very old ($> 10$ Gyr) AGB stars (Wallerstein 2002), or from more massive ($\geq 1 M_\odot$) and younger objects (Willson & Templeton 2009). Recent findings about RVTs also involve binarity, to distinguish their evolutionary channel (Manııck et al. 2018). This means that RVTs are not reliable as distance indicators, and we did not take them into account for our PL relations. Finally, following the classification in Soszyński et al. (2017), we also discarded pWVs, which are a subclass of WVs that likely belong to binary systems (Soszyński et al. 2008). They are overluminous (0.3–0.5 mag in the $K_s$ band Rippeıı et al. 2015) when compared to WVs with similar periods. This means that they do not follow the PL relation of T2Cs and cannot be used as distance indicators. We note that pWVs in Fig. 2 are not more luminous than WVs at the same period. Nonetheless, we did not take them into account for the quoted reasons.

(3) No semi-empirical calibration. In principle, it is not possible, in our case, to adopt a semi-empirical calibration (empirical slope from our own sample and zero-point from literature) for the PLs because our targets are not all located at the same distance. This is true in principle, even though our empirical slope is identical within 0.01 mag to that by Bhardwaj et al. (2017a), as stated in Sect. 4.1.

(4) Selection of the calibration. Keeping these points in mind, we have searched the literature, where several calibrations of the PL relation of T2Cs in the NIR are available (Matsuııga et al. 2006, 2009; Di Criscienzo et al. 2007; Rippeıı et al. 2015; Bhardwaj et al. 2017a; Gaia Collaboration 2017, henceforth PL$_{M06}$, PL$_{D07}$, PL$_{M09}$, PL$_{R15}$, PL$_{B17}$ and PL$_{C17}$). We rule out the preliminary calibration PL$_{C17}$, based on Gaia DR1 parallaxes, because it yields DM$_{L,M,C}$ > 18.9 mag, which does not agree with literature values nor with the Gaia DR1 calibration of RRLs and Classical Cepheids derived in the same work. We note that PL$_{M09}$, PL$_{R15}$ and PL$_{B17}$, which are all based on T2Cs in the LMC, do not provide absolute calibrations of their PLs. Therefore, we adopted the distance of the LMC based on eclipsing binaries (Pietrzyński et al. 2013, DM$_{L,M,C} = 18.493 \pm 0.008 \pm 0.047$ mag) to set the zero-point. In principle, we could calibrate the zero-point using Baade-Wesselink parallaxes for field T2Cs obtained by Feast et al. (2008), but this would be based only on two objects, with a strong difference among the different values of the parallaxes in the literature. After checking each of these calibrating PLs, we decided to adopt PL$_{B17}$. The choice was guided by the following reasons. First, the calibrating PL must be based on both BLHs and WVs. This rules out PL$_{D07}$ since they only used BLHs pulsation models. Second, of the three PLs from the LMC (PL$_{M09}$, PL$_{R15}$, and PL$_{B17}$), the latter is based on light curves with an average of 50 epochs, which is more than twice that of the other two together. Moreover, for targets that are outside their surveyed sky area, they include data from the previous works. Third, we excluded the M06 calibration because it is based on all the three subclasses of T2Cs, including RVTs, which we did not include in our distance analysis. In the end, we adopted PL$_{B17}$, which is based only on BLHs and WVs, which are the subclasses of variables for which we did estimate the distance.

As stated in Sect. 2, we simultaneously estimated distances and the 3D extinction $E(I - K_s)_{B17}$. This was made iteratively, following the method of Bhardwaj et al. (2017b). Normally, after the second or third iteration, the values of distance and reddening converge. We derived individual distances for 710 variables, which are listed in Table 2. Taking all sources of uncertainty into account (uncertainty on the mean magnitude, on the extinction, on the coefficients of the calibrating PL, and the intrinsic width of the PL), the relative uncertainties are in the range 8%–9% for 676 objects, 9%–20% for 38, and only one, at a distance of 2.87 kpc, has a relative uncertainty of $\sim 27\%$.

4.3. Spatial distribution of Type II Cepheids

Knowing the coordinates ($l, b$) and the distances $d$, we derived the coordinates $x_{GAAL}$, $y_{GAAL}$ and $z_{GAAL}$. We adopted a reference frame with the Sun at the origin ($x_{GAAL} = y_{GAAL} = z_{GAAL} = 0$), $x_{GAAL}$ which increases toward the Galactic center, $y_{GAAL}$ which is on the Galactic plane and positive for $l > 0$, and $z_{GAAL}$ which is perpendicular to the Galactic plane and positive for $b > 0$.

Figure 3 clearly shows that several objects are located either in front of or beyond the bulge. The individual distances range from 2.0 to 111.7 kpc, with 11 objects more distant than 30 kpc, most likely belonging to the outer halo, which dominates at Galactocentric distances greater than $\sim 20$ kpc (Carollo et al. 2007, 2018).

The two panels of Fig. 3 also show as black circles with error bars the means and standard deviations of the distances projected onto the Galactic plane (top panel) or onto the longitude ($l = 0$) plane (bottom panel), in angular areas of $2^\circ$. We point out that in the top panel, the most peripheral area at positive $l$ is $4^\circ$ wide, and in the bottom panel, the most central area is $3^\circ$ wide, to take into account the lower density of objects. These mean and standard deviations were derived by fitting a rescaled histogram with a Gaussian. A detailed explanation is provided below.

Galactic ($b = 0$) plane. The central angular areas share basically identical averages. However, the most peripheral areas ($5^\circ < |l| < 9^\circ$), that is, between 0.7 and 1.25 kpc from the Galactic center, show a slight deviation similar to that of the barred distribution of Red Clump stars (Wegg & Gerhard 2013): closer at positive $l$ and more distant at negative $l$. This agrees with our finding that the PL relation does depend on the $l$ coordinate, and gives further evidence of the ellipsoidal symmetry of the old stellar component of the bulge, as has been outlined by Pietrukowicz et al. (2015), who found an ellipsoidal structure for bulge RRLs, with an inclination of the major axis with respect to
the Sun of $20^\circ \pm 3^\circ$, similar to the orientation of the Galactic bar ($\sim 30^\circ$ Gonzalez et al. 2011; Wegg & Gerhard 2013).

**Longitude ($l=0^\circ$) plane.** A distinct yet fictitious trend of the average distance with $b$ is displayed in the bottom panel of Fig. 3. The variables that are closer to the Galactic plane appear to have smaller distances. This is due to a selection effect caused by extinction in the original OGLE catalog, which more easily detects stars in the closest part of the bulge than stars in the farthest part, which are more heavily reddened. The extinction ratio between the $I$ band of OGLE and the $K_s$ band of VVV ranges within a factor of four to ten, and the lower limit of $A_{K_{\text{stellar}}}$ for our targets is $\sim 0.3$ mag. We have considered two other possible explanations for this trend: either an overestimate of reddening at low $b$ or the effect of crowding on aperture photometry, but none can explain the quoted behavior. We discard the possibility that reddening is overestimated because other experiments that adopted either the G12 or the S14 map, which are consistent between themselves, provided estimates of $R_0$, all at about 8.3 kpc (Gonzalez et al. 2012; Bhardwaj et al. 2017b; Majaess et al. 2018). We also rule out the possibility that crowding affects the magnitudes from aperture photometry, making the targets brighter. A direct comparison of PSF versus aperture photometry does not reveal any clear trend with $b$, and the average difference of mean magnitudes is $\Delta K_s(\text{Aperture} - \text{PSF}) = 0.03 \pm 0.013$ mag.

With individual distances and coordinates for our targets, we can estimate the distance of the Galactic center $R_0$. However, the calculation is not straightforward, and cuts and resampling are needed to take the biases into account.

First, we selected only stars at $R_G = \sqrt{r^2_{\text{GAL}} + y^2_{\text{GAL}}}$, which is the distance of the star, projected onto the Galactic plane, between 6 and 11 kpc, to avoid non-bulge stars within the sample. The choice is justified by the results of Pietrukowicz et al. (2015, see their Fig. 5), who showed that the density of bulge stars is very low (lower than $-10\%$ of the peak) outside this distance range. Second, we only selected stars at $b > 3^\circ/0$ and $b < -3^\circ/0$ to avoid the OGLE selection bias. A similar cut was applied by Pietrukowicz et al. (2015), who only used RRLs at $b < -2^\circ/7$ to estimate $R_0$.

These selection criteria left us with 172 stars. However, their average longitude ($\langle l \rangle$) is $0^\circ/275$, which means that the sample is biased toward shorter distances. To overcome this bias, we performed a resampling of the data by randomly selecting 75 stars at negative $l$ and 75 at positive $l$. If $\langle l \rangle$ of the 150 random targets is lower in absolute values than $0^\circ/1$, we kept the sample, otherwise, we repeated the random target selection. We point out that this process is not a proper bootstrap method because we did not allow sampling the same element more than once.

With this set of 150 targets, we plot the distribution of $R_G$ in bins of $0.25$ kpc as shown in Fig. 4 (black histogram). However, this distribution is biased and shifted to greater distances. At fixed coordinates ($l$, $b$), the volume within a given sky area ($\Delta l \Delta b$) and a given depth range ($\Delta d$) increases with distance. This means that the number of stars within the volume (and therefore the probability of detecting a target in the volume) increases quadratically with distance. This causes a bias that shifts the distribution toward the more probable larger distances. To take this geometric effect into account, we scaled the distribution by $d^{-2}$. We fit the scaled distribution (red histogram in Fig. 4) with a Gaussian. We estimated the abscissa of the peak ($x_0$) and adopted it as our estimate of $R_{00}$ on the 8th resampled set. Starting from the random extraction of 150 targets, this process was repeated 5000 times to avoid any selection bias. We adopted an overall average of the 5000 estimates of $R_{00}$ as our final estimate of $R_0$. Based on a sample with $\langle l \rangle =$...
0.016 ± 0.050, we obtain a final $R_0$ estimate of 8.46 kpc. The statistical uncertainty of both $R_0$ was derived as half of the range between the 15.8% and 84.1% percentiles of the distribution of $R_{0,10}$ (see Fig. 5). These thresholds were chosen to enclose 68.3% of the estimates provided by the simulations, like a ±1σ range in a Gaussian distribution. We derived a range of 8.43–8.49 kpc for $R_0$, which means a statistical uncertainty of 0.03 kpc.

We calculated the systematic uncertainty as the squared sum of the average uncertainty on the mean magnitude (0.020 mag), the average uncertainty on the extinction (0.021 mag), and the average propagation of the uncertainties of the calibrating PL coefficients (almost vanishing, 0.001 mag). These together are 0.028 mag, which is 0.11 kpc at 8.30 kpc. Our estimate of $R_0$ does not agree very well with estimates from similar works, either using T2Cs (8.34 ± 0.03 [stat.] ± 0.41 [syst.], Bhardwaj et al. 2017b) or RRLs (8.33 ± 0.05 [stat.] ± 0.14 [syst.], Dékány et al. 2013; 8.27 ± 0.01 [stat.] ± 0.40 [syst.], Pietrukowicz et al. 2015). However, all the quoted papers adopted the reddening law by Nishiyama et al. (2009), which provides a higher $R_{IK} = 0.528$ and, in turn, smaller distances. Had we adopted the Nishiyama et al. (2009) reddening law with our data, it would have provided $R_0 = 8.30 ± 0.03$ (stat.) ± 0.11 (syst.), which would agree perfectly well with the quoted papers. This is evidence of how crucial a correct understanding of the reddening law is. Our estimate of $R_0$ agrees within 1σ with the best overall recommended value from a recent review (~8.3 ± 0.2 [stat.] ± 0.4 [syst.] kpc, de Grijs & Bono 2016).

Finally, a more detailed analysis of the data allows us to show further evidence of the asymmetrical distribution of T2Cs around the Galactic center. By resampling 5000 times only T2Cs at positive $l$, we obtain an average peak of the distribution of 8.29 ± 0.09 kpc. The same process on T2Cs at negative $l$ provides a value of 8.68 ± 0.05 kpc. Together with the distribution of average distances in the Galactic plane (top panel of Fig. 3) and the dependence of the PL relation on $l$, this is strong evidence that T2Cs trace an old, ellipsoidal stellar population.

5. Kinematics
5.1. Proper motion of the center of mass

The proper motion of Sgr $A^*$, the supermassive black hole at the center of the Milky Way, based on VLBA measures, is

$$\mu_{\alpha(Sgr~A^*)} = -6.379 ± 0.026$ mas yr$^{-1}$; $\mu_{\delta(Sgr~A^*)} = -0.202 ± 0.019$ mas yr$^{-1}$ (Reid & Brunthaler 2004). Assuming that the center of mass of the old population traced by T2Cs overlaps with Sgr $A^*$ and has the same proper motion, we adopt the proper motions of T2Cs to obtain an indirect estimate of the proper motion of the center of mass.

As discussed in Sect. 2, we collected proper motions from three different catalogs: VIRAC, PSF, and Gaia. For the analysis in this section, we rejected, from all three catalogs proper motions with a combined statistical error

$$\Delta \mu = \sqrt{\Delta \mu_{\alpha,\text{stat.}}^2 + \Delta \mu_{\delta,\text{stat.}}^2}$$

$$\Delta \mu = \sqrt{\Delta \mu_{\alpha,\text{syst.}}^2 + \Delta \mu_{\delta,\text{syst.}}^2}$$

larger than 2 mas yr$^{-1}$, leaving 553, 343, and 837 targets from VIRAC, PSF, and Gaia, respectively. We point out that the error propagation for the Gaia proper motions, when converting from ($\mu_{\alpha,\text{Gaia}}, \mu_{\delta,\text{Gaia}}$) into ($\mu_{\alpha,\text{PSF}}, \mu_{\delta,\text{PSF}}$), was performed taking into account the covariance terms as suggested by Luri et al. (2018). Gaia coordinates were preprocessed from their native J2015.5 epoch to J2000, the same reference epoch as for PSF, to perform the quoted conversion.

It is crucial to remember that while VIRAC and PSF provide relative proper motions in the frame of reference of the Galaxy, Gaia provides absolute proper motions in the practically inertial frame of reference defined by quasars. This allows an interesting comparison among the catalogs. As a first step, we left out VIRAC proper motions. As displayed in Figure 6, the distribution of $\Delta \mu_{\alpha} = \mu_{\alpha,\text{Gaia}} - \mu_{\alpha,\text{PSF}}$ is centered at $\Delta \mu_{\alpha,\text{peak}} = -6.41 ± 0.02$ mas yr$^{-1}$. For the $b$ component, we find $\Delta \mu_{\delta,\text{peak}} = 0.12 ± 0.03$ mas yr$^{-1}$. These numbers were derived using 251 T2Cs for which we have both Gaia and PSF proper motions and that are located within 2 kpc from the center of the Galaxy, as derived in Sect. 4.3. The latter criterion was adopted as a compromise to leave out possible thick-disk stars and to retain a large sample of targets.

When we assume that the velocities of our targets are randomly distributed around the center of mass, which is reasonable because this is the behavior of old, low-mass bulge stars (Spenhauer et al. 1992; Minniti 1996; Babusiaux et al. 2010; Kunder et al. 2016), then ($\Delta \mu_{\alpha,\text{peak}}, \Delta \mu_{\beta,\text{peak}}$) is an indirect estimate of the proper motion of the center of mass of the Galaxy. This is supported by the fact that $\Delta \mu_{\alpha,\text{peak}}$ is identical well within 1σ to the longitudinal component of the proper motion of Sgr $A^*$. The latitudinal component $\Delta \mu_{\delta,\text{peak}}$ does not agree with that of Sgr $A^*$, but the mean uncertainty on the proper motions (0.82 mas yr$^{-1}$ for PSF and 0.47 for Gaia) is larger than the offset. We also checked whether there is a trend of $\mu_{\alpha}$ and $\mu_{\delta}$ with distance, and found nothing significant in either the Gaia or the PSF sample. This implies that there is no evidence of net rotation.

We have performed the same analysis using proper motions from VIRAC and Gaia and found that in this case, $\Delta \mu_{\delta,\text{peak}} = -5.43 ± 0.03$ mas yr$^{-1}$ and $\Delta \mu_{\beta,\text{peak}} = -0.04 ± 0.02$ mas yr$^{-1}$. While the $b$ component is similar to that of Sgr $A^*$, the component along $l$ is different from $\mu_{\alpha(Sgr~A^*)}$ by as much as ~1 mas yr$^{-1}$. Moreover, we found that for 265 targets in common between the VIRAC and PSF datasets, the difference between the medians of $\mu_{\alpha}$ and $\mu_{\beta}$ are ~0.97 ± 2.62 and +0.12 ± 2.62 mas yr$^{-1}$, respectively. No dependence of $\mu_{\alpha}$ on $l$ or on $b$ was found in either of the two catalogs. We conclude that the $\mu_{\beta,\text{PSF}}$ are more reliable than $\mu_{\beta,\text{VIRAC}}$ not only a priori, as discussed in Sect. 2, but also a posteriori. Because these are both relative measures of the proper motion, the assumption of randomly distributed motions around the Galactic center would imply a zero median $\mu_{\beta}$. The comparisons with Gaia and that with the
PSF sources both indicate that VIRAC proper motions are systematically shifted by \( \sim 1 \) mas yr\(^{-1} \) to the east on the Galactic plane.

### 5.2. Velocity dispersion

Tangential velocities can be used to test the kinematic properties of a stellar population. If bulge T2Cs trace a spheroidal population that is kinematically hot, then their distribution should be dominated by a velocity dispersion with negligible rotation such as the Bulge RRL population (Minniti 1996; Kunder et al. 2016; Marconi & Minniti 2018; Contreras Ramos et al. 2018; and references therein). In this case, the velocity ellipsoid should be fairly symmetric in the Galactic longitude \( (\mu_l, v_l) \) and latitude \( (\mu_b, v_b) \) components. To test this hypothesis, we derived the tangential velocity \( (\mu_b, v_b) \) and its two components \( v_b \) and \( \mu_b \), using the classical relation \( v_l = 4.74 \cdot d \cdot \mu \), where \( d \) is in kpc and \( \mu \) in mas yr\(^{-1} \). We used \( d \) as derived in Sect. 4.2 and \( \mu \) from \( \text{Gaia} \) because it is the most complete sample with the smallest uncertainties. While pWVs and RVTs might have valid values of \( \mu \), we cannot derive their \( v_b \), since \( d \) is not available. Table 3 displays the proper motions and velocities.

Finally, we derived the standard deviation of \( v_b \) and \( \mu_b \) and obtained 106.2 \( \pm \) 3.0 km s\(^{-1} \) and 97.3 \( \pm \) 5.5 km s\(^{-1} \), respectively. These values were corrected by subtracting, in quadrature, the average uncertainties on \( v_b \) (20.5 km s\(^{-1} \)) and \( \mu_b \) (9.9 km s\(^{-1} \)). After this correction, we find \( \sigma_{v_b} = 104.2 \pm 3.0 \) km s\(^{-1} \) and \( \sigma_{\mu_b} = 96.8 \pm 5.5 \) km s\(^{-1} \). The agreement between the two is better than 1.2\( \sigma \), thus providing further evidence that the T2Cs belong to a kinematically hot population. We cannot rule out the possibility that despite our cuts, we still included some thick-disk objects. Stricter selections concerning the target distance from the center are hampered by the sample size.

### 6. Conclusions

We have retrieved \( K_s \)-band light curves from VVV aperture photometry for 894 of 924 T2Cs in the OGLE IV catalog (Soszynski et al. 2017). We calculated mean magnitudes and amplitudes based on PLOESS fits (Persson et al. 2004; Braga et al. 2018) to the light curves. For BLHs and WVs, we simultaneously estimated individual extinctions and distance moduli, based on a 3D reddening map (Schultheis et al. 2014) and on a PL relation. The calibration of the PL relation was based on the slope and zero-point of T2Cs in the LMC (Bhardwaj et al. 2017a), anchored with a late-type eclipsing binaries distance to the LMC (Pietrzyński et al. 2013). We found distances ranging from 2.0 to 111.7 kpc, which means that our objects are located in the bulge, in the inner and outer halo, and possibly in the thick disk. The mean individual relative uncertainty is 8.6\% independent of distance and with a small standard deviation of 1.2\%.

The distribution of the individual distances, taking various geometric and selection biases into account, provides an estimate of the distance of the Galactic center \( R_0 \) of 8.46 \( \pm \) 0.03(stat.) \( \pm \) 0.11(syst.), which agrees with the recommended value of 8.3 \( \pm \) 0.2(stat.) \( \pm \) 0.4(syst.) kpc.

![Fig. 6. Left: distribution of \( \Delta \mu_b \) for 251 T2Cs located less than 2 kpc away from the center of the Galaxy. A Gaussian fit to the distribution is shown in red. The abscissa of the peak and the \( \sigma \) of the fit are labeled. Right: same as left panel, but for \( \Delta \mu_a \).](A51_page_9.png)

Table 3. Proper motions and velocities of target T2Cs.

| ID | \( \mu_v \) mas yr\(^{-1} \) | \( \mu_a \) mas yr\(^{-1} \) | \( \mu_b \) mas yr\(^{-1} \) | \( \mu_l \) mas yr\(^{-1} \) | \( \mu_r \) mas yr\(^{-1} \) | \( v^l \) km s\(^{-1} \) | \( v^b \) km s\(^{-1} \) | \( v_t^l \) km s\(^{-1} \) |
|----|-----------------|-----------------|-----------------|-----------------|-----------------|---------------|---------------|---------------|
| 001 | ...             | ...             | ...             | ...             | ...             | ...           | ...           | ...           |
| 002 | \( -3.13 \pm 0.81 \) | \( 1.75 \pm 0.04 \) | ...             | ...             | ...             | ...           | ...           | ...           |
| 003 | \( -3.46 \pm 0.81 \) | \( -1.05 \pm 0.06 \) | ...             | ...             | ...             | ...           | ...           | ...           |
| 004 | ...             | ...             | ...             | ...             | ...             | ...           | ...           | ...           |
| 005 | \( 2.63 \pm 1.48 \) | \( 1.68 \pm 0.06 \) | ...             | ...             | ...             | ...           | ...           | ...           |
| 006 | ...             | ...             | ...             | ...             | ...             | ...           | ...           | ...           |
| 007 | ...             | ...             | ...             | ...             | ...             | ...           | ...           | ...           |
| 008 | \( 0.71 \pm 0.92 \) | \( -1.25 \pm 0.12 \) | ...             | ...             | ...             | ...           | ...           | ...           |
| 009 | ...             | ...             | ...             | ...             | ...             | ...           | ...           | ...           |
| 010 | ...             | ...             | ...             | ...             | ...             | ...           | ...           | ...           |

Notes. Only the first 10 of the 924 lines of the table are shown. The full table is available at the CDS. The velocities are based on the absolute proper motions by \( \text{Gaia} \).
Our estimate of $R_0$ does not agree with other estimates with similar methods ($R_0 \approx 8.30$ kpc, Dékány et al. 2013; Pietrukowicz et al. 2015; Bhardwaj et al. 2017b), but the difference is consistent with the different reddening law that was adopted (Alonso-García et al. 2017 instead of Nishiyama et al. 2009).

We provided solid evidence that the old stellar population in the bulge is ellipsoidal. First, we found a non-negligible dependence of the PL relation on the $l$ coordinate. This has been described before by Groenewegen et al. (2008), but their limited sample hampered the precision of the coefficient ($-0.028 \pm 0.031$ mag), while ours is more precise ($-0.019 \pm 0.003$ mag). Second, we found that at $l \leq -5^\circ$, the average distance is larger, while at $l \geq 5^\circ$ the average distance of T2Cs is smaller, on a map projected onto the Galactic plane. Third, which is a similar but more quantitative approach as the second point, we found that the distribution of T2Cs at positive $l$ is centered at 8.29 ± 0.09 kpc, while that of T2Cs at negative $l$ is centered at 8.68 ± 0.05 kpc.

We also adopted proper motions from both Gaia and VVV itself to constrain the kinematic properties of T2Cs in the bulge. The analysis was restricted to only the sources with a combined statistical error smaller than 2 mas yr$^{-1}$. The power of the synergy between Gaia and VVV astrometric data is clear when comparing the absolute proper motions from Gaia with relative proper motions from the VVV. The mean difference ($-6.41 \pm 0.02$ and 0.12 ± 0.03 mas yr$^{-1}$ in the longitude and latitude direction, respectively) for T2Cs within 2 kpc from the Galactic center is similar within the uncertainties (0.82 mas yr$^{-1}$ for PSF and 0.47 mas yr$^{-1}$ for Gaia) to the VLBA estimate of the relative proper motion of Sgr A* ($-6.379 \pm 0.026$ and $-0.202 \pm 0.019$ mas yr$^{-1}$). This is reasonable if we assume that the T2Cs of the bulge belong to the kinematically hot, old stellar population (Minniti 1996; Kunder et al. 2016). Another piece of evidence supporting both this assumption is that the velocity dispersion in both the longitude and latitude directions agree within almost 1$\sigma$ ($\sigma_v = 104.2 \pm 3.0$ km s$^{-1}$, $\sigma_v = 96.8 \pm 5.5$ km s$^{-1}$). The difference may be due to contamination by thick-disk stars in the 2 kpc sphere around the Galactic center.

It is important to note that while the distribution and kinematics of metal-rich populations in the bulge, tracing the X-shaped structure, have been studied widely, the distributions of more metal-poor populations based on different tracers remain to be investigated in detail. This work on T2Cs provides results that are consistent with RRLs. The spectroscopic follow-up of these objects in the near future will allow us to confirm the differences in their spatial distributions and kinematics to those of metal-rich populations in the Galactic bulge.

**Note added in proof.** Recent investigations based on the position of T2Cs in the Magellanic Clouds, showed that WV's might not be exclusively old stars, but rather a mix of intermediate-age and old stars Iwanek et al. 2018. We have checked that the tilted distribution found in Sect. 4.3 does not substantially change if traced only by BLHs (purely old) and by WV's, although the tilt is more clear when traced by the latter.

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