Primeval very low-mass stars and brown dwarfs – IV. New L subdwarfs, Gaia astrometry, population properties, and a blue brown dwarf binary

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ABSTRACT
We present 27 new L subdwarfs and classify five of them as esdL and 22 as sdL. Our L subdwarf candidates were selected with the UKIRT Infrared Deep Sky Survey and Sloan Digital Sky Survey. Spectroscopic follow-up was carried out primarily with the OSIRIS spectrograph on the Gran Telescopio Canarias. Some of these new objects were followed up with the X-shooter instrument on the Very Large Telescope. We studied the photometric properties of the population of known L subdwarfs using colour–spectral type diagrams and colour–colour diagrams, by comparison with L dwarfs and main-sequence stars, and identified new colour spaces for L subdwarf selection/study in current and future surveys. We further discussed the brown dwarf transition-zone and the observational stellar/substellar boundary. We found that about one-third of 66 known L subdwarfs are substellar objects, with two-thirds being very low-mass stars. We also present the Hertzsprung–Russell diagrams, spectral type–absolute magnitude corrections, and tangential velocities of 20 known L subdwarfs observed by the Gaia astrometry satellite. One of our L subdwarf candidates, ULAS J233227.03+123452.0, is a mildly metal-poor spectroscopic binary brown dwarf: a ∼L6p dwarf and a ∼T4p dwarf. This binary is likely a thick disc member according to its kinematics.

Key words: brown dwarfs – stars: chemically peculiar – stars: Population II – stars: substellar objects – binaries: spectroscopic

1 INTRODUCTION

Very low-mass stars (VLMS) and brown dwarfs (BDs) are classified using M, L, T, and Y types according to spectral morphology that is dominated by temperature-dependent chemistry and thermal properties (Bessell 1991; Kirkpatrick, Henry, & McCarthy 1991; Jones et al. 1994; Tsuji et al. 1996;
Kirkpatrick et al. 1999a, 2012; Martín et al. 1999; Burgasser et al. 2002, 2003b; Cushing et al. 2011). The spectral types of field VLMS extend from M to early-type L (Dieterich et al. 2014; Dupuy & Liu 2017). Most massive BDs have spectral types between late-type M and mid-type L depending on their age (Zhang et al. 2017b). For 0.1–1 Gyr massive BDs are late-type M, evolving to later type as they cool (L type for ~1–10 Gyr age, and T type for ages >10 Gyr; e.g., fig. 8 of Burrows et al. 2001).

The classification of VLMS and BDs with subsolar metallicity is complicated by the population’s wide metallicity range, which leads to significant diversity in the strength of spectral features. And available sample size also places some limitations on the scope of classification (Gizis 1997; Kirkpatrick 2005; Lépine, Rich, & Shara 2007; Jao et al. 2008; Dhillon et al. 2012; Burgasser, Cruz, & Kirkpatrick 2007; Kirkpatrick et al. 2014; Zhang et al. 2017a). However, previous work has established a set of subdwarf classes that are indicative of metallicity. Gizis (1997) refined M classification by setting up three subclasses; dwarf (dM), subdwarf (sdM) and extreme subdwarf (esdM). An additional intermediate sub-class d/sd for late-type M and L subdwarfs was also suggested by Burgasser, Cruz, & Kirkpatrick (2007), aimed at intermediate metallicity to further encompasses the Galactic thick disc population. Lépine, Rich, & Shara (2007) revised the M dwarf classification into four subclasses; dM, sdM, esdM, and ultra-subdwarf (usdM).

Kirkpatrick (2005) proposed a three-parameter classification strategy for late-type M, L, and T dwarfs to indicate their metallicity, temperature/clouds, and gravity. For example, a metal-poor halo object 2MASS J05325346+8246465 (Burgasser et al. 2003a) was classified esdL7 and an intermediate metal-poor object SDSS J141624.12+134827.4 (SD1416; Bowler, Liu, & Dupuy 2010; Cunningham et al. 2010; Schmidt et al. 2010) was classified sdL7 in Kirkpatrick et al. (2010). And more recently Zhang et al. (2017a) classifies L subdwarfs using three subclasses: sdL, esdL, and usdL, based on the relative strength of subsolar metallicity sensitive spectral features across the optical and near-infrared (NIR). The metallicity ranges of the usdL, esdL, and sdL subclasses are approximately [Fe/H] ≤ −1.7; −1.7 ≤ [Fe/H] ≤ −1.0; and −1.0 ≤ [Fe/H] ≤ −0.3, respectively (see Zhang et al. 2017a).

The L subdwarf population is composed of metal-deficient low-mass stars and high-mass BDs with effective temperature (T_{eff}) in the range of ~1300–2700 K. They have strong FeH absorption bands, weak or absent VO and CO bands, and enhanced collision-induced H$_2$ absorption (CIA H$_2$; Bates 1952; Saumon et al. 2012), as summarized in Zhang et al. (2017a). Sample identification has been enabled by modern large-scale optical and NIR surveys: the Two Micron All Sky Survey (2MASS; Skrutskie et al. 2006), the Sloan Digital Sky Survey (SDSS; York et al. 2000), the UKIRT Infrared Deep Sky Survey (UKIDSS; Lawrence et al. 2007), and the Wide-field Infrared Survey Explorer (WISE; Wright et al. 2010), with a total of 39 L subdwarfs previously reported in the literature (Bowler, Liu, & Dupuy 2010; Burgasser et al. 2003a, 2004a; Burgasser et al. 2004b; Burgasser & Kirkpatrick 2006; Cunningham et al. 2010; Schmidt et al. 2010; Cushing et al. 2009; Gizis & Harvin 2006; Kirkpatrick et al. 2010, 2014, 2016; Lépine et al. 2002; Luhman & Sheppard 2014; Lodieu et al. 2010, 2012, 2017; Scholz, Lodieu, & McCaughrean 2004a; Scholz et al. 2004b; Schneider et al. 2016; Sivarani et al. 2009; Smith et al. 2018; Zhang et al. 2017a,b, 2018).

In general, subdwarfs with sdM–sdL subclass are kinematically associated with the Galactic thick disc, while those with esdM–esdL and usdM–usdL subclasses have kinematics of the Galactic halo (Burgasser et al. 2008b; Cushing et al. 2009; Zhang et al. 2013, 2017a). By comparison to the population of known L dwarfs, L subdwarfs in the solar neighbourhood are rare. This is partly because they are predominantly thick disc and halo objects (cf. Reddy, Lambert, & Allende Prieto 2006, have shown that the fractions of thick disc and halo stars in the solar neighbourhood are 7 and 0.6 per cent), but also because L subdwarfs occupy a narrower mass range than L dwarfs (Zhang et al. 2017b). Lodieu et al. (2017) report a UKIDSS/SDSS surface density (for late M and early L dwarfs) of around 0.04 per deg$^2$, pointing to potential for a much larger (hundreds strong) detectable population across surveyed sky. Inroads into this population are important for the detailed study of observed diversity and population make-up amongst the L subdwarfs. This is the fourth paper in a series titled Primeval very low-mass stars and brown dwarfs. The first paper reported the discovery of six new L subdwarfs, defined a new L subdwarf classification scheme, and studied the atmospheric properties of ultra-cool subdwarfs based on 22 late-type M and L subdwarfs (Zhang et al. 2017a, hereafter Paper I). In the second paper, we presented the most metal-poor substellar object, and a procedure to distinguish massive halo BDs from the least-massive stars. We also found that mid-type L to early-type T subdwarfs of the Galactic halo are located in a substellar subdwarf gap, known also as the halo BD transition-zone. This zone covers a narrow mass range but spans a wide $T_{eff}$ range due to unsteady nuclear fusion (Zhang et al. 2017b, hereafter Paper II). In the third paper, we presented the discovery of three new halo transitional BDs (Zhang et al. 2018, hereafter Paper III). Here, we present the discovery of 27 new L subdwarfs and their population properties, as well as a spectroscopic blue BD binary. Observations are presented in Section 2. Classification is carried out in Section 3. Section 4 assesses L subdwarf population properties. Section 5 presents Gaia observations of L subdwarfs. Section 6 sums up and presents our conclusions.

2 OBSERVATION

2.1 Candidate selection

The L subdwarf candidates in our programme were generally selected from the UKIDSS Large Area Survey (LAS) and the SDSS, following the selection criteria/procedure described in Paper I. In addition two extra candidates without SDSS detections were selected from a ULAS proper motion catalogue, covering 1500 deg$^2$ to an approximate 5σ depth of $J = 19.6$ (Smith et al. 2014). These two objects (ULAS J130710.22+151130.4 and ULAS J144151.55+043738.5) both have relatively blue $J$–$K$ colour and proper motion higher than 0.4 arcsec per year.

Our programmatic candidate sample consisted of 64 objects, of which five had been confirmed as L subdwarfs by
| Name               | SpT   | SDSS i | SDSS z | Y (MKO) | J (MKO) | H (MKO) | K (MKO) | Ref⁹ |
|--------------------|-------|--------|--------|---------|---------|---------|---------|------|
| ULAS J133348.27+273505.5 | sdL1  | 20.50 ±0.05 | 19.20 ±0.06 | 18.48 ±0.05 | 17.93 ±0.05 | 18.06 ±0.11 | 18.08 ±0.17 | —   |
| ULAS J135204.39+014202.7 | sdL0  | 19.49 ±0.02 | 18.01 ±0.02 | 16.98 ±0.01 | 15.16 ±0.01 | 14.56 ±0.01 | 14.39 ±0.01 | —   |
| ULAS J135216.31+312327.0 | sdL1  | 21.01 ±0.04 | 18.66 ±0.05 | 17.69 ±0.02 | 16.93 ±0.02 | 16.66 ±0.03 | 16.41 ±0.04 | —   |
| ULAS J141319.18+005009.4 | sdL0  | 20.48 ±0.04 | 18.66 ±0.04 | 17.62 ±0.02 | 16.82 ±0.02 | 16.40 ±0.02 | 15.93 ±0.02 | —   |
| ULAS J231912.16+045126.1 | sdL1  | 20.78 ±0.05 | 19.03 ±0.05 | 17.69 ±0.02 | 16.87 ±0.02 | 16.47 ±0.02 | 15.60 ±0.03 | —   |
| VVV J133847.27+37505.5 | sdL1  | 20.52 ±0.05 | 18.76 ±0.04 | 17.47 ±0.02 | 16.62 ±0.02 | 15.98 ±0.02 | 15.28 ±0.02 | —   |
| ULAS J133836.97+292201.7 | sdL0  | 22.53 ±0.28 | 20.10 ±0.15 | 18.56 ±0.06 | 17.37 ±0.03 | 16.81 ±0.04 | 16.37 ±0.05 | —   |
| ULAS J134206.86+053742.9 | sdL0  | 21.83 ±0.11 | 19.45 ±0.09 | 18.46 ±0.04 | 17.54 ±0.03 | 17.21 ±0.04 | 16.78 ±0.05 | —   |
| ULAS J134420.75+063513.4 | sdL0  | 22.41 ±0.27 | 20.42 ±0.17 | 19.12 ±0.08 | 18.46 ±0.10 | 18.08 ±0.12 | 18.04 ±0.20 | —   |
| ULAS J134749.27+333359.8 | sdL1  | 21.97 ±0.03 | 18.07 ±0.02 | 16.66 ±0.01 | 15.85 ±0.01 | 15.46 ±0.01 | 15.27 ±0.02 | —   |
| ULAS J134852.93+101611.8 | sdL0  | 20.76 ±0.05 | 19.03 ±0.04 | 18.46 ±0.04 | 17.37 ±0.03 | 16.81 ±0.04 | 16.37 ±0.05 | —   |
| ULAS J135359.58+101801.1 | sdL1  | 21.49 ±0.14 | 19.03 ±0.04 | 18.26 ±0.04 | 17.36 ±0.03 | 16.89 ±0.04 | 15.92 ±0.04 | —   |
| WISE J155011.89+085707.2 | sdL0  | 21.29 ±0.09 | 19.14 ±0.05 | 17.54 ±0.02 | 16.33 ±0.01 | 15.95 ±0.01 | 15.43 ±0.02 | —   |
| SDSS J141001.13+223812.5 | sdL0  | 18.58 ±0.01 | 15.91 ±0.01 | 14.26 ±0.01 | 12.99 ±0.01 | 12.47 ±0.01 | 12.05 ±0.01 | —   |
| SDSS J141832.35+052332.0 | sdL0  | 20.11 ±0.04 | 18.27 ±0.03 | 16.86 ±0.01 | 16.00 ±0.01 | 15.61 ±0.01 | 15.19 ±0.01 | —   |
| SDSS J141851.55+043738.7 | sdL0  | 18.45 ±0.04 | 17.23 ±0.03 | 16.95 ±0.04 | 16.34 ±0.04 | 16.04 ±0.04 | 15.64 ±0.04 | —   |
| ULAS J151649.84+083607.1 | sdL1  | 22.65 ±0.24 | 20.25 ±0.10 | 18.74 ±0.04 | 17.35 ±0.02 | 16.71 ±0.03 | 16.26 ±0.03 | —   |
| ULAS J154638.34+011213.0 | sdL1  | 22.10 ±0.14 | 19.98 ±0.11 | 18.56 ±0.05 | 17.51 ±0.04 | 17.21 ±0.05 | 16.95 ±0.08 | —   |
| WISE J175610.80+281528.2 | sdL1  | —       | —       | —       | —       | 14.66 ±0.11 | 14.19 ±0.11 | 13.79 ±0.10 | —   |
| LSR J182611.3+301419.1 | sdL0  | —       | —       | —       | —       | 11.61 ±0.11 | 11.22 ±0.10 | 10.78 ±0.13 | —   |
is an sdL3 subdwarf according to a new optical-NIR spectrum we obtained with the X-shooter.

We have also obtained optical-NIR spectroscopy of five of our candidates, a known sdL8 subdwarf and three radial velocities. The subclasses of L subdwarfs are based on their metallicity and compared optical spectra of some candidates to those of L dwarfs. To assign spectral type/class to our candidates, we combined the spectra of these objects from both epochs.

2.2 GTC spectroscopy

Candidates have been followed up with optical spectroscopy using the System for Imaging for low-Resolution Integrated Spectroscopy (OSIRIS; Cepa et al. 2000) instrument on the Gran Telescopio Canarias (GTC) since 2012. Table 2 shows a summary of our GTC spectroscopic observations. The spectra were mostly obtained using the R500R grism, with three making use of the R300R grism. These provide resolving power of around 500 and 300, respectively. The spectra were reduced using standard procedures with IRAF. The HgAr, Ne, and Xe arcs were used on the wavelength calibration of our spectra. Spectral flux calibrations were achieved with standard stars (Bohlin, Colina, & Finley 1995; Hamuy et al. 1994; Oke 1974, 1990). Spectral features caused by telluric absorptions in these spectra were not corrected. Five objects were observed twice to gain additional exposure time, and for the purposes of this paper we combined the spectra of these objects from both epochs.

2.3 VLT spectroscopy

We have also obtained optical-NIR spectroscopy of five of our candidates, a known sdL8 subdwarf and three radial velocities. L dwarfs using the X-shooter spectrograph (Vernet et al. 2011) on the Very Large Telescope (VLT) since 2014. Table 3 shows a summary of our VLT spectroscopic observations. These X-shooter spectra were observed in an ABBA nodding mode with a 1.2 arcsec slit providing a resolving power of 6700 in the VIS arm and 4000 in the NIR arm. These spectra were reduced with ESO Reflex in a lazy mode (Freudling et al. 2013). The flat field used for the science data was also used for flat-fielding the flux standard star spectrum (by set flat strategy to false). Telluric corrections were achieved for both VIS and NIR arms using telluric standards that were observed right after or before our targets at a very close airmass. To extract telluric spectra for correction, we fitted and normalized spectra of telluric standards to their continuum and removed non-telluric features (e.g. hydrogen absorption lines) with IRAF SPLIT. These X-shooter spectra were smoothed by 101 and 51 pixels in the VIS and NIR arms, respectively, for display in figures, which reduced the resolving power to about 600 and increased their signal-to-noise ratio (SNR per pixel) by 10 and 7 times.

3 CLASSIFICATION

In Paper I, we classified L subdwarfs into sdL, esdL, and usdL subclasses. The subtypes of L subdwarfs are based on the comparison of their optical spectra to those of L dwarfs. The subclasses of L subdwarfs are based on their metallicity sensitive spectral features in the optical and NIR.

3.1 Optical classification

To assign spectral type/class to our candidates, we compared their optical spectra to those of established L subdwarf standards and identified the ones that provided the closest match to the overall spectral profile and to known spectral features sensitive to low metallicity. We also compared optical spectra of some candidates to those of L dwarf standards, when we could not find a close match to available L subdwarf standards. We used TIO and VO absorption bands in the optical region to distinguish L subdwarfs from L dwarfs, taking into account the standard/model spectral comparisons shown in

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1 IRAF is distributed by the National Optical Astronomy Observatory, which is operated by the Association of Universities for Research in Astronomy, Inc., under contract with the National Science Foundation.

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Table 1. Continued.

| Name                | SpT  | SDSS i | SDSS z | Y (MKO) | J (MKO) | H (MKO) | K (MKO) | Ref |
|---------------------|------|--------|--------|---------|---------|---------|---------|-----|
| ULAS J223440.80+001002.6 | sdL3 | 22.05±0.14 | 20.16±0.12 | 18.97±0.07 | 17.63±0.04 | 17.15±0.05 | 16.90±0.07 | 26 |
| ULAS J225902.14+115602.1 | sdL0 | 21.30±0.16 | 19.23±0.11 | 17.97±0.02 | 17.05±0.02 | 16.63±0.03 | 16.21±0.03 | 26 |
| ULAS J230256.53+121310.2 | sdL0 | 21.74±0.15 | 19.46±0.08 | 18.30±0.04 | 17.49±0.03 | 17.05±0.06 | 16.73±0.05 | 26 |
| ULAS J230443.30+093423.9 | sdL0 | 21.65±0.16 | 19.27±0.08 | 18.21±0.03 | 17.23±0.02 | 16.74±0.03 | 16.27±0.03 | 26 |

Notes. a ATLAS i, z and VHS Y JHK photometry. b ATLAS i, z and VIKING Y JHK photometry. c ATLAS i, z and VHS Y JHK photometry. The other NIR photometry of objects with Y band detection are from UKIDSS. J, H, and K photometry of the other objects without UKIDSS Y band detection are converted from 2MASS with equations (1–3). e VVV J12564163–6202039 is an sdL3 subdwarf according to a new optical-NIR spectrum we obtained with the X-shooter. f Reference: 1. Burgasser et al. (2004a); 2. Burgasser et al. (2004a); 3. Burgasser (2004b); 4. Burgasser & Kirkpatrick (2006); 5. Burgasser, Cruz, & Kirkpatrick (2007); 6. Burgasser et al. (2009); 7. Bowler, Liu, & Dupuy (2010); Burningham et al. (2010); Schmidt et al. (2010); 8. Cushing et al. (2009); 9. Gizis & Harvin (2006); 10. Kirkpatrick et al. (2010); 11. Kirkpatrick et al. (2014); 12. Kirkpatrick et al. (2016); 13. Lépine et al. (2002); 14. Luhan & Sheppard (2014); 15. Lodieu et al. (2010); 16. Lodieu et al. (2012); 17. Lodieu et al. (2017); 18. Scholz, Lodieu, & McCaughrean (2004a); 19. Scholz et al. (2004b); 20. Schneider et al. (2016); 21. Sivarani et al. (2009); 22. Smith et al. (2018); 23. Paper I; 24. Paper II; 25. Paper III; 26. This work.
New L subdwarfs and population properties

Fig. 1 of Paper I. This figure shows that the TiO absorption bands at around 710, 775, and 850 nm are getting stronger from the dL to sdL subclasses (from [Fe/H] = 0.0 to −0.5), and then getting gradually weaker from the sdL to usdL subclasses (from [Fe/H] = −0.5 to −2.0). The strong TiO absorption bands are a signature of early-type sdL, which is not the case for early type L dwarfs. The VO absorption band at 788–810 nm gets weaker from early-type dL to sdL subclasses and disappears in the esdL subclass ([Fe/H] < −1.0). Furthermore, a characteristic difference between early-type sdL and early-type esdL subclasses can be seen in the 770–810 nm range, where early-type esdL spectra follow a sloping straight line but early-type sdL spectra follow a dipping curve due to stronger TiO and weaker VO absorption.

Fig. 1 shows six of our new L subdwarfs that have stronger metal-poor spectral features with very weak or absent VO absorption bands at around 800 nm. Five of these have been classified as esdL, with one (UL0212+06) classified as sdL (but close to the sdL/esdL border). The VO absorption band strength of these subdwarfs is similar to the synthetic spectral models with [Fe/H] ≤ −1.0. By comparison, Fig. 2 shows 20 of our new L subdwarfs whose spectral features are indicative of slightly higher metallicity than those in Fig. 1. The spectra of these objects compared well with known sdL subdwarfs or showed typical sdL spectral features (strong TiO absorption bands).

3.1.1 Six L subdwarfs with strong metal-poor features

ULAS $J111429.54+072809.5$ (UL1114+07) was classified esdL0, since it compares well with ULAS J033350.84+001406.1 (UL0333; Lodieu et al. 2012), an esdL0 classified in Paper I.

ULAS $J135216.31+312327.0$ (UL1352+31), ULAS $J145234.65+043738.4$ (UL1452+04), and ULAS $J223302.03+062030.8$ (UL2233+06) were all classified esdL0.5 since they compare well with ULAS J124425.90+102441.9 (UL1244; Lodieu et al. 2012), an esdL0.5 classified in Paper I. We also note that UL2233+06 compared best to UL1244 at 600–840 nm. And UL1352+31 has slightly more flux than UL1244 at 770–810 nm, which is indicative of a slightly lower metallicity. Since UL1244 has [Fe/H] ≈ −1.5 (Paper I), UL1352+31 is close to the esdL/usdL boundary.

ULAS $J021258.08+064115.9$ (UL0212+06) was classified sdL1 since it compares well with the sdL1 subdwarf SDSS J133348.24+273508.8 (SD1333; Paper I). It may have a flatter plateau over 0.738–0.757 $\mu$m than SD1333, though this could be due to noise or telluric effects. The flux of UL0212+06 beyond 880 nm is not well calibrated due to the lack of second-order correction for the OSIRIS spectrum. Although UL0212+06 (and SD1333) are classified as sdL they have weaker 800 nm VO absorption than any other sdL dwarfs, and thus lie close to the sdL/esdL boundary.

ULAS $J231924.35+052524.5$ (UL2319+05) has a similar spectral profile to SD1333, but has weaker TiO absorption bands and slightly more flux over 770–810 nm (with a sloping straight line morphology) indicative of a lower metallicity. We therefore classified this object as esdL1.
Figure 2. Optical spectra of 21 new L subdwarfs (black) compared to known L subdwarfs (red) and dwarfs (magenta). Spectra are normalized near 840 nm. Spectra of SD1347, 2M0041, and UL0216 are from Paper I. Spectra of 2M1439, 2M0355, 2M1155, and 2M1632 are from Kirkpatrick et al. (1999a). The spectrum of 2M1756 is from Kirkpatrick et al. (2010) and was corrected for telluric absorption. The spectrum of 2M1507 is from Reid et al. (2000). 2M0645 was observed by our X-shooter follow-up programme.
3.1.2 Eighteen sdL0–8 subdwarfs

Thirteen of our objects in Fig. 2 compare well with the sdL0 subdwarf SDSS J134749.74+333601.7 (SD1347; Paper I) and were thus classified sdL0. These objects are ULAS J011840.73+084424.7, ULAS J023803.12+054526.1, ULAS J075335.23+200622.4, ULAS J082206.61+044101.8, ULAS J123142.99+015045.4, ULAS J124104.75−000531.4, ULAS J125226.62+092920.1, ULAS J134852.93+101611.8, ULAS J135359.58+011856.7, ULAS J141832.35+025323.0, ULAS J225902.14+115602.1, ULAS J230256.53+121310.2, and ULAS J230443.30+093423.9.

ULAS J134206.86+053724.9 (UL1342+05) was classified sdL0.5, since it compares well with 2MASS J00412179+3547133 (2M0041; Burgasser et al. 2004a), an sdL0.5 classified in Paper I.

ULAS J223440.80+001002.6 (UL2234+00) was classified sdL1, since it compares well to the sdL1 subdwarf 2M23 17561080+2815238 (2M1756; Kirkpatrick et al. 2010), and shows stronger TiO absorption at around 710, 780 and 850 nm than the L1 dwarf standard 2MASS J14392836+1929149 (2M1439; Kirkpatrick et al. 1999a).

ULAS J134423.98+280603.8 (UL1344+28) and ULAS J14151.55+043738.5 (UL1415+04) compared well to the
L4 standard 2MASS J1155009+230706 (2M1155; Kirkpatrick et al. 1999a). However, both have somewhat stronger TiO absorption at 850 nm (than this L4 dwarf) that is a signature of the sdL subclass. UL1344+28 and UL1441+04 compared well with the sdL4 subdwarf ULAS J021642.97+004005.6 (UL0216; Paper I) and were thus classified sdL4. We note that UL1344+28 has slightly weaker 850 nm TiO absorption than UL0216, and thus a slightly higher metallicity than UL0216. UL1441+04 was previously reported as a blue L1 dwarf by an independent search for L and T dwarfs (Marocco et al. 2015).

**ULAS J130710.22+151105.4 (UL1307+15)** compares well to the L8 dwarf standard 2MASS J16322911+1904407 (2M1632; Kirkpatrick et al. 1999a) at wavelengths below 860 nm. The metal-poor sensitivity of stronger TiO absorption bands below 860 nm become insignificant in the spectra of late-type sdL subclasses. However, metal hydride (CrH, FeH) absorption for late-type sdL subdwarfs is stronger than that in late-type L dwarfs. UL1307+15 has stronger 870 nm FeH absorption than 2M1632, indicating a lower metallicity. UL1307+15 compares well with the sdL8 subdwarf 2MASS J06453153–6646120 (2M0645; Kirkpatrick et al. 2010), particularly in the red optical (including the 870 nm FeH band), and it was thus classified sdL8.

### 3.1.3 Three sdL subdwarfs with new subtypes

Three objects did not closely compare with any of the L subdwarf or L dwarf standards, but showed clear evidence of L subdwarf nature.

**ULAS J154348.34−011213.0 (UL1546−01)** compares reasonably to the optical spectral profile of the L3 dwarf standard 2MASS J03554191+2257016 (2M0355; Kirkpatrick et al. 1999a), but has stronger TiO absorption at around 710 and 850 nm. Therefore, we classified this object as sdL3.

**ULAS J141203.85+121609.9 (UL1412+12)** and **ULAS J151649.84+083607.1 (UL1516+08)** both compare reasonably to the L5 dwarf standard 2MASS J15074769−1627386 (2M1507; Reid et al. 2000), but have stronger TiO absorption at around 850 nm. We therefore classified these objects as sdL5. UL1516+08 has a slightly weaker 850 nm TiO band compared to UL1412+12, and likely has somewhat higher metallicity. These objects were both previously reported as blue L4 and L5 dwarfs respectively, by Marocco et al. (2015).

We propose UL1412+12 as an sdL5 standard, since we have obtained good-quality optical and NIR spectra with the GTC and VLT (Section 3.2). Note that a previously classified sdL5 object, 2MASS J06164006−6407194 (2M0616; Cushing et al. 2009), was re-classified as esdL6 in Paper I. Also, ULAS J153508.85+081506.8 (Lodieu et al. 2010) was originally classified as sdL5 but was re-classified as sdL3.5−4 in Lodieu et al. (2017) and as usdL3 in Paper I and III. And a suspected sdL5 subdwarf, WISEA J135501.90−825838.9 (WI1355; Kirkpatrick et al. 2010), seems less secure as a spectroscopic standard, since its $J−H$ and $J−K$ colours are similar to those of L5 dwarfs.

From Table 1, we can see that we still missing some spectral subtypes of L subdwarfs. These missing subtypes including sdL2, sdL6, sdL9; esdL2, esdL5, esdL8−9; and usdL2 and usdL5−9. Most of them are late L subtypes as late L dwarfs are fainter and more difficult to identify with current facilities.

### 3.2 NIR constraints

L dwarfs/subdwarfs have rather diverse NIR spectral features because they are distributed into a wide range of metallicity and gravity. Objects that are more metal-poor have lower opacity, thus have more flux at shorter wavelength, and tend to be more compact (i.e. higher gravity). NIR spectral fluxes are suppressed gradually as metallicity decreases due to enhanced CIA H$_2$ (see fig. 9 in Paper I). The optical-NIR morphology of L subdwarfs is thus very characteristic. And within the NIR itself, it is relatively easier to distinguish sdL from dL subdwarf using NIR spectral features (which are more significant than optical features such as TiO absorptions for sdL subclasses). NIR spectra are also very useful for distinguishing usdL from esdL subclasses because the spectral variations caused by the metallicity differences are larger in the NIR than in the optical.

UL2233+06 was observed under our X-shooter follow-up programme. The spectrum of this object has SNR of $\sim$8 at 820 nm and $\sim$6 at 1300 nm, and was smoothed by 101 (VIS) and 51 (NIR) pixels for display. Fig. 3 shows the optical to NIR spectrum of UL2233+06 compared (via optical normalization) to those of the sdL0.5 subdwarf 2MASS J00412179+3547133 (2M0041; Burgasser et al. 2004a; Paper I) and the L0 dwarf 2MASP J0345432+254023 (SD0345; Kirkpatrick et al. 1999a). UL2233+06 has much stronger NIR flux suppression than the sdL0.5 subdwarf, and its CO absorption is absent. This is entirely consistent with the esdL0.5 classification of UL2233+06, and the dL–sdL–esdL sequence clearly shows the gradually increasing level of NIR suppression.

Note that Fig. 3 does not show the true relative flux of these subclasses. Fig. 16 in Paper I shows that esdL and dL subclasses with close spectral type have similar $H$ band absolute magnitudes. Thus relative flux is best shown by normalising spectra in the $H$ band, as in Fig. 4. This figure shows the same three objects from Fig. 3 but with the additional L0$_\gamma$ dwarf 2MASS J01415823−4633574 (2M0141; Kirkpatrick et al. 2006). Although these objects all have spectral subtype L0–L0.5 (classified in optical) they have very different physical parameters. UL2233+06 is older with lower metallicity, but relatively higher mass and Teff. While at the other extreme the L0$_\gamma$ dwarf is young with higher metallicity, cooler Teff, and much lower mass and gravity (Filippazzo et al. 2015; Faherty et al. 2016).

UL1412+12 was also observed by our X-shooter follow-up programme. The original spectrum has SNR of $\sim$6 at 910 nm and $\sim$30 at 1300 nm, and was smoothed by 101 (VIS) and 51 (NIR) pixels for display. Fig. 5 shows that UL1412+12 has largely suppressed NIR flux and a weaker CO absorption band compared to the L5 dwarf standard SDSS J083506.16+195304.4 (SD0835; Chiu et al. 2006; Kirkpatrick et al. 2010). This is as expected for the sdL5 classification (from the optical).

We obtained the optical to NIR spectra of UL1307+15 and 2M0645 under our X-shooter follow-up programme. The original spectrum of UL1307+15 has SNR of $\sim$4 at 900 nm and $\sim$7 at 1300 nm. The original spectrum of 2M0645 has SNR of $\sim$20 at 900 nm and $\sim$45 at 1300 nm. Fig. 6 shows...
Table 2. Summary of the characteristics of the spectrophotometric observations made with the OIIRIS on the GTC. R500R and R300R grisms cover a wavelength range of 480-1020 nm. A 0.8 arcsec slit used for all observations.

| ULAS Name | Proposal ID | UT date | Seeing("") | Airmass | Grism | T_int (s) | SpT | Std (SpT)Ref. |
|-----------|-------------|---------|------------|---------|-------|----------|-----|--------------|
| J111429.54+072809.5 | GTC39-12B | 2013-01-17 | 0.67 | 1.24 | R500R | 900 × 1 | esdL0 | GD 140 (DA2.2)3 |
| J135216.31+312327.0 | GTC4-14A | 2014-03-12 | 0.80 | 1.15 | R500R | 600 × 1 | esdL0.5 | GD 153 (DA1.2)3 |
| J145234.65+043738.4 | GTC39-12B | 2013-01-20 | 1.00 | 1.14 | R500R | 900 × 1 | esdL0.5 | G191-B2B (DA8)3 |
| J223302.03+062030.8 | GTC39-12B | 2012-09-02 | 0.70 | 1.11 | R500R | 1800 × 1 | esdL0 | G191-B2B (DA8)3 |
| J231924.35+052524.5 | GTC39-12B | 2012-09-02 | 0.60 | 1.13 | R500R | 900 × 1 | esdL1 | G191-B2B (DA8)3 |
| J011840.73+084424.7 | GTC63-13A | 2013-08-01 | 0.80 | 1.09 | R500R | 1200 × 3 | sdL0 | Ross 640 (DAZ5.5)3 |
| J021258.08+064115.9 | GTC39-12B | 2012-09-01 | 0.70 | 1.08 | R500R | 1800 × 1 | sdL0.5 | Ross 640 (DAZ5.5)3 |
| J023803.12+054526.1 | GTC60-15A | 2015-08-23 | 0.70 | 1.18 | R300R | 900 × 1 | sdL0 | Ross 640 (DAZ5.5)3 |
| J075353.25+200622.4 | GTC39-12B | 2012-12-06 | 0.75 | 1.15 | R500R | 900 × 1 | sdL0 | G191-B2B (DA8)3 |
| J082206.61+044101.8 | GTC39-12B | 2012-12-06 | 0.75 | 1.11 | R500R | 900 × 1 | sdL0.5 | G191-B2B (DA8)3 |
| J123142.99+015045.4 | GTC60-15A | 2015-03-16 | 0.90 | 1.52 | R500R | 900 × 1 | sdL0 | Hilt 600 (B1)2 |
| J124114.75+005351.4 | GTC39-12B | 2015-03-08 | 0.70 | 1.68 | R500R | 1200 × 3 | sdL0 | GD 190 (DB2) |
| J125226.62+092920.1 | GTC39-12B | 2015-03-16 | 0.70 | 1.21 | R500R | 900 × 1 | sdL0.5 | Hilt 600 (B1)2 |
| J130710.22+151103.4 | GTC4-14A | 2014-07-26 | 0.50 | 1.80 | R500R | 1200 × 3 | sdL8 | GD 153 (DA1.2)3 |
| J134206.86+053724.9 | GTC4-14A | 2014-07-25 | 0.80 | 1.50 | R500R | 1200 × 2 | sdL1 | GD 153 (DA1.2)3 |
| J134423.98+280630.8 | GTC63-13A | 2013-04-08 | 0.90 | 1.06 | R500R | 1200 × 3 | sdL4 | GD 190 (DB2)3 |
| J134852.93+101118.8 | GTC63-13A | 2013-03-19 | 0.90 | 1.37 | R500R | 900 × 1 | sdL0 | L 1363-3 (DQ6)3 |
| J135539.58+011856.7 | GTC4-14A | 2014-07-28 | 0.70 | 1.73 | R500R | 900 × 3 | sdL0 | GD 153 (DA1.2)3 |
| J141203.85+121609.9 | GTC39-12B | 2013-01-20 | 1.00 | 1.09 | R500R | 1800 × 1 | sdL5 | G191-B2B (DA8)3 |
| J141832.35+025323.9 | GTC39-12B | 2014-01-17 | 0.67 | 1.12 | R500R | 900 × 1 | sdL0 | GD 140 (DA2.2)3 |
| J141415.55+043738.5 | GTC4-14A | 2014-07-19 | 0.60 | 1.34 | R500R | 1200 × 3 | sdL4 | GD 153 (DA1.2)3 |
| J151649.84+083607.1 | GTC4-14A | 2014-07-25 | 0.70 | 1.36 | R500R | 1200 × 3 | sdL5 | GD 153 (DA1.2)3 |
| J134638.34+112113.0 | GTC63-13A | 2014-05-07 | 0.90 | 1.17 | R500R | 1200 × 3 | sdL3 | L 1363-3 (DQ6)3 |
| J223440.80+001002.6 | GTC63-13A | 2014-07-17 | 0.80 | 1.15 | R500R | 1200 × 3 | sdL1 | L 1363-3 (DQ6)3 |
| J225902.14+115602.1 | GTC63-13A | 2014-07-18 | 0.95 | 1.18 | R500R | 900 × 2 | sdL0 | GD 140 (DA2.2)3 |
| J230443.30+093423.9 | GTC39-12B | 2014-07-20 | 0.60 | 1.06 | R500R | 1200 × 2 | — | GD 153 (DA1.2)3 |
| J230256.53+121310.2 | GTC63-13A | 2013-07-18 | 0.80 | 1.08 | R500R | 1550 × 2 | sdL0 | GD 190 (DB2)3 |
| J230443.30+093423.9 | GTC63-13A | 2013-07-13 | 0.70 | 1.08 | R500R | 1550 × 2 | sdL0 | LDS 748B (DB4)3 |

Notes. References for flux calibration standard stars are 1. Bohlin, Colina, & Finley (1995); 2. Hamuy et al. (1994); 3. Oke (1974); 4. Oke (1990).
that UL1307+15 has a similar spectral profile to 2M0645 from optical to H band. We note that UL1307+15 has more K band flux and slightly narrower H band spectral profile than 2M0645. UL1307+15 is worth of further investigation (e.g. on binarity).

To confirm the subdwarf status of the other 19 L subdwarfs (Fig. 2) without NIR spectra, we instead compared their optical to NIR photometric spectral energy distributions (SEDs) to those of known L dwarfs and subdwarfs with the same optical spectral subtypes (Fig. 7). These 19 L subdwarfs all show significantly suppressed NIR SEDs compared to those of the L dwarfs.

The 13 new L0 subdwarfs have similar SEDs to LSR J182611.3+301419.1 (LSR1826; Lépine et al. 2002), which was classified as sdL0 in Paper I. UL1342+05 shows a similar SED to the sdL0.5 2M0041. Although there is no sdL3 subdwarf to compare with UL1546−01, it can be seen that the UL1546−01 SED is more suppressed relative to the L3 dwarfs, and by a level similar to the suppression amongst the sdL0es (when compared to the L0 dwarfs). UL1344+28 and UL1441+04 have similar SEDs to the sdL4 UL0216, UL1412+12 has a slightly stronger NIR suppression than UL1516+08 indicating a slightly lower metallicity. And this is also supported by its slightly stronger 850 nm TiO absorption band when compared to UL1516+08.

### 3.3 Astrometry and radial velocity

One of our new L subdwarfs (UL0753+20) was observed by the Gaia Collaboration et al. (2018) astrometric survey. Eleven of our new L subdwarfs (including UL10753+20) were observed by the UKIDSS second epoch survey, and are therefore included in the UKIDSS proper motion catalogue of Smith et al. (2014). Proper motions of the rest of our objects were measured from SDSS, UKIDSS and PS1 epochs following the procedure described in Zhang et al. (2009). We estimated spectroscopic distances for these 27 new L subdwarfs based on the relationship between spectral type and H band absolute magnitude, which is less sensitive to metallicity/subclass for L subdwarfs. The relationship was derived from 20 L subdwarfs observed by the Gaia mission (Gaia Collaboration et al. 2016, see Section 5). Table 4 shows the proper motions, distances and tangential velocities of these 27 new L subdwarfs. They are distributed at distances from ~ 50 to 250 pc. Seventeen of them have tangential velocities higher than 100 km s$^{-1}$. Seven of them have tangential velocities above 200 km s$^{-1}$, which includes these six L subdwarfs with the strongest metal-poor features shown in Fig. 1.

We note that UL0212 (the sdL in Figure 1 that is close to the sdL/esdL border) has tangential velocity of 329±71 km s$^{-1}$, which is similar to the space motions of the five esdL subdwarfs in Fig. 1. Thus the most metal-poor objects of the sdL subclass (UL0212 and SD1333) both have typical halo kinematics.

We measured the radial velocities (RV) of UL2233+06, UL1412+12, UL1307+15, and 2M0645 based on their X-shooter spectra. To measure the RVs of L subdwarfs in our X-shooter follow-up programme, we also observed three RV standard L dwarfs [DENIS-P J144137.3−094559 (Martin et al. 1999), an L0.5 dwarf with an RV of −27.9 ± 1.2 km s$^{-1}$ (Bailer-Jones 2004); 2MASS J08354256−0819237 (Cruz et al. 2003), an L5 dwarf with an RV of 29.89 ± 0.06 km s$^{-1}$ (Blake, Charbonneau, & White 2010); and DENIS-P J025503.3−470049 (Martín et al. 1999), an L9 dwarf with an RV of 17.5 ± 2.8 km s$^{-1}$ (Zapatero Osorio et al. 2007)]. All X-shooter spectra were smoothed by 21 and 11 pixels in the VIS and NIR respectively, to increase their SNR. We first measured the RV differences between our objects and the RV standards with the closest spectral subtype using cross-correlation on their strong absorption lines (Na I and Cs I in the optical and K I in the NIR). Measured RV differences were then corrected for barycentric effects, and converted into final RV using the known RVs of the standards. The final RVs of UL2233+06, UL1412+12, UL1307+15, and 2M0645 are −164 ± 15, −57 ± 12, −60±18, and −33 ± 10 km s$^{-1}$, respectively.

The esdL+usdL and the sdL subclasses are approximately kinematically associated with the halo population and thick disc populations, respectively. The ratio between the number of esdL+usdL and sdL subdwarfs in our UKIDSS-SDSS sample is much higher than that between halo and thick disc stars measured by Reddy, Lambert, & Allende Prieto (2006). This is likely due to our selection bias, as esdL or usdL subdwarfs can more easily be picked out by their extreme colours. There are likely a lot more sdL subdwarfs observed by existing photometric surveys, but not yet confirmed by spectroscopy. This arises because sdL subdwarf identification is more problematic than esdL and usdL, due to contamination from scattered M and L dwarfs in colour–colour diagrams. There are 66 L subdwarfs known to-date, including 7 usdL, 20 esdL, and 39 sdL subdwarfs.

### Table 3. Summary of the characteristics of the spectroscopic observations made with X-shooter. Wavelength ranges for the VIS and NIR arms are 530–1020 and 990–2480 nm. 1.2 arcsec slits are used for both VIS and NIR arms. DENIS-P J025503.3−470049, 2MASS J08354256−0819237 and DENIS-P J144137.3−094559 were observed as RV standards.

| Name | SpT | UT date | Seeing | Air | T$_{int}$(VIS) | T$_{int}$(NIR) | Telluric (SpT) | Airm |
|------|-----|---------|--------|-----|--------------|--------------|----------------|------|
| ULAS J024035.36+060629.3 | sdM7 | 2015-01-17 | 0.97" | 1.42 | 12×237 s | 12×250 s | HD 22686 (A0 V) | 1.27 |
| ULAS J141203.85+121609.9 | sdL5 | 2015-02-25 | 0.96" | 1.26 | 12×287 s | 12×300 s | HIP 76069 (B9 V) | 1.33 |
| ULAS J233227.03+123452.0 | L6p+T4p | 2015-09-10 | 1.03" | 1.28 | 12×285 s | 12×300 s | HIP 117927 (B9V) | 1.20 |
| ULAS J223302.03+060203.8 | esdL0.5 | 2015-09-11 | 1.20" | 1.24 | 12×285 s | 12×300 s | HIP 105164 (B7 III) | 1.09 |
| ULAS J130710.22+151031.4 | sdL8 | 2018-04-23 | 1.03" | 1.34 | 12×285 s | 12×300 s | HIP 61257 (B9.5 V) | 1.15 |
| 2MASS J06453153−3661612 | sdL8 | 2016-02-19 | 1.09" | 1.35 | 8×140 s | 8×150 s | HD 77281 (A3 IV) | 1.26 |
| DENIS-P J025503.3−470049 | L9 | 2015-08-12 | 1.57" | 1.44 | 8×140 s | 8×150 s | HIP 105164 (B7 III) | 1.36 |
| 2MASS J08354256−0819237 | L5 | 2015-10-31 | 0.62" | 1.31 | 8×140 | 8×150 | HIP 38734 (B9 V) | 1.33 |
| DENIS-P J144137.3−094559 | L0.5 | 2016-03-19 | 0.66" | 1.11 | 8×140 | 8×150 | HIP 76836 (B9.5 V) | 1.04 |

**MNRA 000, 1–77 (2018)**
New L subdwarfs and population properties

Figure 7. Optical–NIR photometric spectral energy distributions of 19 new L subdwarfs (black lines and blue diamonds) and an L1 dwarf (UL2247+05; magenta line and cyan hexagons), compared to those of L dwarfs (grey lines and green crosses) and known L subdwarfs (red lines and magenta open circles). The spectrum of LSR1826 is from Burgasser et al. (2004a). The spectrum of 2M1756 is from Kirkpatrick et al. (2010). The spectrum of UL0216 is from Paper I. Spex spectra of L0–L5 dwarfs used for comparison are of 2MASS J13313310+3407583 (L0); 2MASS J01340281+0508125 (L0; Kirkpatrick et al. 2010); 2MASSI J2107316−030733 (L0), SDSS J202820.32+005226.5 (L3), 2MASSI J11040281+0508125 (L4; Burgasser et al. 2004a); 2MASSW J0228110+253738 (L0), 2MASS J12122770+0257198 (L0), 2MASSW J0208183+254253 (L1), SDSS J104842.84+011158.5 (L1; Burgasser et al. 2008a); 2MASS J20343769+0827009 (L1), 2MASS J08234818+2428577 (L3), 2MASSW J1146345+223053 (L3), 2MASS J22425317+2542573 (L3), 2MASS J13571237+1428398 (L4), 2MASS J13571237+1428398 (L4), 2MASS J17461199+5034036 (L5; Burgasser et al. 2010); SDSSp J053951.99−005902.0 [L5; M. Cushing (unpublished); Fan et al. 2000]; SD0835 (L5; Chiu et al. 2006).
Meanwhile, there are a few thousand photometric candidates (e.g. Skrzypek, Warren, & Faherty 2016) and spectroscopically confirmed L dwarfs to-date (e.g. Best et al. 2018). The ratio between known esdL+usdL subdwarfs and L dwarfs is comparable to the fraction of halo population in the solar neighbourhood (e.g. 0.6%; Reddy, Lambert, & Allende Prieto 2006).

3.4 A blue BD binary

**ULAS J233227.03+123452.0** (UL2332+12) was selected as one of our ultra-cool subdwarf candidates. This object has previously been identified by Skrzypek, Warren, & Faherty (2016), who flagged it as a candidate binary from SED fitting. Their subsequent spectroscopic assessment classified it as a T1.5p, and following a more detailed NIR spectroscopic analysis they concluded it was more likely to be a metal-poor T0 dwarf. We re-assess the nature of this object based on our full optical-NIR spectroscopy.

Fig. 8 shows that the optical spectrum of UL2332+12 compares well with the L7 dwarf standard 2MASS J08505953+1057156 (2M0850; Kirkpatrick et al. 1999a) below 860 nm, but has relatively more flux at longer wavelength. In the red optical UL2332+12 compares better to the sdL7 subdwarf SD1416. However, UL2332+12 and SD1416 are quite different in the NIR.

UL2332+12 was observed by our X-shooter follow-up programme, and Fig. 9 shows the full optical–NIR spectrum.

### Table 4. Proper motions, distances and tangential velocities of our 27 new L subdwarfs.

| Name       | $\mu_{\alpha}$ (mas/yr) | $\mu_{\delta}$ (mas/yr) | Distance (pc) | $V_{tan}$ (km s$^{-1}$) |
|------------|--------------------------|--------------------------|---------------|--------------------------|
| UL0118+03  | 14±2                     | −80±12                   | 209±20        | 80±18                    |
| UL0212+06  | 2±4                      | −425±9                   | 139±13        | 282±27                   |
| UL0238+05  | 186±18                   | −145±15                  | 95±9          | 106±15                   |
| UL0753+20  | −36±1                    | −191±1                   | 74±7          | 68±7                     |
| UL0822+04  | 35±7                     | −154±5                   | 92±9          | 69±8                     |
| UL1114+07  | −17±9                    | −306±6                   | 211±18        | 306±21                   |
| UL1231+01  | −225±6                   | 22±4                     | 164±18        | 176±18                   |
| UL1241+00  | −61±9                    | −45±5                    | 245±21        | 89±14                    |
| UL1252+09  | −299±7                   | 6±6                      | 117±10        | 166±15                   |
| UL1307+15  | −391±14                  | −124±13                  | 86±8          | 168±17                   |
| UL1342+05  | −106±8                   | −212±5                   | 142±12        | 160±15                   |
| UL1344+28  | −345±11                  | 120±10                   | 88±5          | 153±15                   |
| UL1348+10  | −283±11                  | −179±9                   | 109±10        | 174±18                   |
| UL1352+31  | 542±8                    | −111±29                  | 148±13        | 388±22                   |
| UL1353+01  | −303±12                  | −37±10                   | 142±12        | 67±7                     |
| UL1412+12  | −227±10                  | −32±10                   | 54±5          | 58±7                     |
| UL1418+02  | −3±7                     | −277±6                   | 79±8          | 103±10                   |
| UL1441+04  | −264±8                   | −508±8                   | 98±9          | 266±24                   |
| UL1452+04  | −46±9                    | −291±11                  | 157±15        | 220±23                   |
| UL1516+08  | −171±8                   | 44±5                     | 79±8          | 66±7                     |
| UL1546+01  | −45±7                    | −115±6                   | 122±12        | 71±9                     |
| UL2233+06  | −320±60                  | 8±50                     | 236±22        | 370±36                   |
| UL2234+00  | −34±9                    | −96±7                    | 145±7         | 70±11                    |
| UL2259+11  | 34±19                    | −115±13                  | 126±12        | 72±15                    |
| UL2302+12  | −123±7                   | −191±6                   | 153±15        | 165±17                   |
| UL2304+09  | −21±13                   | −46±4                    | 133±13        | 32±9                     |
| UL2319+05  | 509±11                   | −401±11                  | 167±14        | 513±50                   |

### Table 5. Properties of a spectral binary UL2332+12. Note the spectral type is from synthesized fit.

| Parameter | UL2332+12 |
|-----------|-----------|
| $\alpha$ (J2000) | 23$^h$32$^m$27$^s$03 |
| $\delta$ (J2000) | +12$^\circ$34'32"0 |
| Epoch | 2009-09-04 |
| SDSS $i$ | 22.64±0.27 |
| SDSS z | 19.91±0.11 |
| PS1 $i$ | 22.00±0.17 |
| PS1 z | 20.23±0.04 |
| PS1 y | 19.13±0.03 |
| UKIDSS Y | 18.10±0.04 |
| UKIDSS J | 16.90±0.02 |
| UKIDSS H | 16.39±0.03 |
| UKIDSS K | 15.88±0.03 |
| WISE W1 | 15.18±0.04 |
| WISE W2 | 14.77±0.07 |
| Spectral type | L6p+T4p |
| Distance (pc) | 58±12 |
| $\mu_{\alpha}$ (mas/yr) | 420±10 |
| $\mu_{\delta}$ (mas/yr) | −107±13 |
| $V_{tan}$ (km s$^{-1}$) | 118±21 |
| RV (km s$^{-1}$) | −35±8 |
| $U$ (km s$^{-1}$) | −85±18 |
| $V$ (km s$^{-1}$) | −86±14 |
| $W$ (km s$^{-1}$) | −28±13 |

Figure 8. Optical spectra of UL2332+12 observed with GTC and VLT compared to those of SD1416 (from SDSS) and 2M0850 (from Kirkpatrick et al. 1999a). Spectra are normalized at 910 nm.
The original spectrum of UL2332+12 has SNR of ∼20 at 900 nm and ∼45 at 1300 nm, and was smoothed by 101 (VIS) and 51 (NIR) pixels for display. UL2332+12 has a similar NIR spectral profile to the sdL7 subdwarf SD1416, but has deeper H$_2$O absorption bands (a feature characteristic of later T dwarfs). UL2332+12 has a similar NIR spectrum to the T2 standard SDSS J075840.33+324723.4 (SD0758; Knapp et al. 2004); however, it has more flux at around 850 nm and no pronounced CH$_4$ absorption. Indeed, the full optical–NIR spectrum of UL2332+12 does not compare very well to any single L dwarf/subdwarf or T dwarf, but does have much more in common with the spectrum of DENIS-P J225210.73−173013.4 (DE2252; Kendall et al. 2004; Reid et al. 2006), an L4+T3.5 spectroscopic binary. L dwarfs with unresolved T dwarf companions can have bluer NIR colours (e.g. fig. 14 of Zhang et al. 2010) and peculiar spectral features compared to normal L dwarfs (Burgasser et al. 2010; Bardalez Gagliuffi et al. 2014; Marocco et al. 2015; Manjavacas et al. 2016).

We compared the NIR spectrum of UL2332+12 to those of synthesized spectral binaries constructed using L and T spectra from the Spex library, as described in Burgasser et al. (2010). Fig. 10 shows best-fitting synthesized spectrum, which combines the L dwarf SDSS J133148.92−011651.4 (SD1331; Hawley et al. 2002) with the T4 dwarf (2MASSI J2254188+312349; Burgasser et al. 2004a). SD1331 is optically classified L6 (Hawley et al. 2002) but also has a NIR classification of L8p (Knapp et al. 2004), a pecu-
lilar blue L dwarf. This synthesized combination is a much better match than the best-fitting single object (2MASS J11582077+0435014(sdL7); Kirkpatrick et al. 2010). We assessed our 78 best-fitting synthesized spectra to constrain spectral type estimates for primary (L6.1±0.5) and the secondary (T4.0±1.7) components. And we note that these best-fitting combinations all include SD1331 (or other blue L dwarfs) as the primary.

Fig. 11 compares the optical–NIR spectrum of SD1331 to that of other optically classified L6 dwarfs, demonstrating the ‘blue’ peculiarity of this object (see also Marocco et al. 2015). By direct comparison UL2332+12 has bluer \( J-K \) and redder \( i-J \) colours than SD1331, consistent with an additional unresolved mid-T companion. In colour-space SD1331 lies \( \sim \)mid-way between the dL and sdL sequences (e.g., it is indicated by a magenta open pentagon in Fig. 17, a figure that will be fully discussed in Section 4.3), and is on the BD side of the stellar-substellar boundary. Objects like SD1331 are thus likely to be mildly metal-poor BDs, and our best-fitting explanation for UL2332+12 is that it is a mildly metal-poor L6+T4 blue BD binary. Note that there is not mildly metal-poor mid-T dwarf in the Spex library available for our synthesized spectroscopic fitting. However, this does not affect our conclusion, as the affect of slightly lower metallicity on the shape and flux ratio of \( J \) and \( H \) band spectra of T dwarfs is negligible (e.g. Pinfield et al. 2012). UL2332+12 has slightly less \( K \) band flux than the synthesized spectrum. This indicates that the T type companion of UL2332+12 is slightly more metal-poor than the T4 dwarf used in the synthesized spectrum.

The \( H \) band absolute magnitudes (MKO) of L6 and T4 dwarfs are around 12.90 and 14.10 mag according to Dupuy & Liu (2012). The combined \( H \) band absolute magnitude of an L6+T4 binary is thus around 12.59 mag, suggesting a distance of \( 58^{+10}_{-13} \) pc for UL2332+12 as a binary. A mildly metal-poor L6 dwarf could have a slightly brighter (~ 0.2 mag) \( H \) band absolute magnitude compared to a normal L6 dwarf, which would lead to a slightly increased distance. However, such an increase (~5 pc) is much less than our estimated distance uncertainty and is thus not significant.

We measured the proper motion and RV of UL2332+12 following the procedure described in Section 3.3. UL2332+12 has proper motion \( \mu_{RA} = 420 \pm 10 \) mas/yr and \( \mu_{Dec} = -107 \pm 13 \) mas/yr, which is measured from its UKIDSS and PS1 epochs with a baseline of 2.74 yr. The tangential velocity of UL2332+12 is thus \( 118^{+35}_{-26} \) km s\(^{-1}\). The RV of UL2332+12 was measured from its X-shooter spectrum and an L9 type RV standard (DENIS-P J025503.3−470049), yielding an RV of \( -35 \pm 8 \) km s\(^{-1}\). The space motion of UL2332+12 is thus \( U = -85 \pm 18, \ V = -86 \pm 14 \) and \( W = -28 \pm 13 \) km s\(^{-1}\). We calculated the halo membership probabilities of UL2332+12 based on kinematics and stellar population fractions of the thin disc (0.93), thick disc (0.07) and halo (0.006) in the solar neighbourhood (Reddy, Lambert, & Allende Prieto 2006). The thin disc and thick disc membership probabilities of UL2332+12 are 22 and 77 per cent, respectively.

A summary of the measured and best-fitting properties of UL2332+12 are presented in Table 5.

### 3.5 M subdwarfs and M–L dwarfs amongst our candidates

About one-third (22) of our candidate L subdwarfs were identified as M6–L1 dwarfs or M7–M9 subdwarfs through spectroscopic follow-up. Table 6 shows spectral types and photometry for these objects.

ULAS J224749.77+053207.9 (UL2247+05) has \( i-J \) and \( J-K \) colours similar to those of mid-type sdL. However, its optical spectrum compares well with the L1 dwarf 2M1349 (see Fig. 12). Its \( i_{P1} - J \) and \( J-K \) colours are consistent with early-type L dwarfs, and we note that its SDSS \( i \) band magnitude has a much larger uncertainty than the PS1 \( i_{P1} \) band magnitude. Fig. 7 shows that UL2247+05 has a slightly bluer SEDs than other L1 dwarfs. UL2247+05 has a proper motion of \( \mu_{RA} = -56 \pm 14 \) mas/yr and \( \mu_{Dec} = -56 \pm 15 \) mas/yr (measured using the method described in Section 3.3). Its distance is \( 146^{+30}_{-25} \) pc based on spectral type and \( H \) band magnitude, leading to a tangential velocity of \( 54^{+18}_{-17} \) km s\(^{-1}\).

When considering the M subdwarfs, there are several classification schemes for the late-types. In the most recent sdM scheme (Lépine, Rich, & Shara 2007), the metallicity consistency of the latest sdM subclasses has not been fully tested across all subtypes, due to the lack of wide binaries containing both early- and late-type M subdwarfs. We have also noticed that the strength of the metallicity features are usually under estimated for late-type M subdwarfs compared to early-type M subdwarfs (section 4.1 of Paper I). As a result, late-type M stars with the same metallicity as early-type esdM subdwarfs can be classified as sdM, and some late-type M stars with the same metallicity as early-type sdM subdwarfs can be classified as M dwarfs. Zhang et al. (2013) noticed that \( \sim \) 18 per cent of SDSS M dwarfs classified on the scheme of Lépine, Rich, & Shara (2007) have thick disc or halo like kinematics (\( V < -100 \) km s\(^{-1}\)). Furthermore, different subclass names are used for objects with similar metallicity in different classification schemes. The late-type d/sdM subclass of Burgasser, Cruz, & Kirkpatrick (2007)
New L subdwarfs and population properties

Figure 13. Optical spectra of newly discovered M subdwarfs (left panel) and dwarfs (middle and right panels) compared to SDSS M dwarf template spectra (Bochanski et al. 2007). The VO absorption band at around 800 nm is highlighted with a blue band in the left-hand panel. Spectra are normalized at 825 nm.

would be expected to have a similar metallicity range as the early-type sdM subclass of Lépine, Rich, & Shara (2007).

We classified our M subdwarfs using their optical metal-poor spectral features, such as weak VO absorption near 800 nm. This represents an extension of our L subdwarf scheme (Paper I) into the late-type sdMs, and allows us to consistently carry out sdM/sdL classification of our programmatic sample. Fig. 13 shows the spectra of these 21 late-type M stars compared to stacked SDSS M6–M9 spectral templates (Bochanski et al. 2007). Seven of these objects show weaker VO absorption bands than the M dwarf templates: ULAS J135122.15+141914.9 has the strongest metal-poor features with the absence of 800 nm VO absorption, and was classified as esdM7. ULAS J002009.35+160451.2, ULAS J020628.22+020255.6, and ULAS J024035.36+060629.3 (UL0240+06) have a weaker 800 nm VO absorption bands than the M dwarf templates, and were classified as sdM9, sdM7, and sdM7, respectively.
ULAS J010756.85+100811.3, ULAS J143154.18−004114.3 and ULAS J143517.18−04173.1 have weaker 800 nm VO absorption bands than the M dwarf templates, but slightly stronger than UL0240+06. We classified these objects as sdM subdwarfs, although they could be classified as dM or d/sdM using existing classification schemes. The other 14 objects were classified as late-type M dwarfs since they compare well with M dwarf templates.

Fig. 14 shows the X-shooter spectrum of UL0240+06. The original spectrum of UL0240+06 has SNR of ~2 at 910 nm and ~4 at 1300 nm, and was smoothed by 101 (VIS) and 51 (NIR) pixels for display. UL0240+06 shows suppressed NIR flux compared to the M7 dwarf VB 8 (Kirkpatrick, Henry, & McCarthy 1991). UL0240+06 has a similar spectral profile to the d/sdM7 object 2MASS J13593574+3031039 (2M1359; Burgasser et al. 2004a).

### 4 L SUBDWARF PHOTOMETRY AND STELLAR/SUBSTELLAR MIX

The label ‘subdwarf’ originally acknowledged the location of these objects below the stellar main-sequence in the Hertzsprung-Russell diagram (HRD), due to their lower opacity and bluer colour resulting from subsolar metallicity. However, the three metallicity subclasses of L subdwarfs exhibit a range of different characteristics, sometimes showing quite substantial differences in spectral energy distribution from their L dwarf counterparts (see figs 14 and 15 of Paper I). In this section we explore the optical to infrared colours of these objects below the main-sequence in the HRD.

#### 4.1 A photometric sample of L subdwarfs

To study the photometric diversity of L subdwarfs, we joined our UKIDSS-SDSS sample with L subdwarfs collected from the literature to form a larger sample of 66 L subdwarfs. We gathered their photometry: SDSS $i$, $z$, UKIDSS (MKO) $Y$, $J$, $H$, $K$. We also used $i$ and $z$ photometry from the VLT Survey Telescope’s (VST) ATLAS survey (Shanks et al. 2015) for a few objects that are not covered by SDSS. We used NIR photometry from the Visible and Infrared Survey Telescope for Astronomy (VISTA) surveys for a few

| Name                  | SpT    | SDSS $i$ | SDSS $z$ | UKIDSS $Y$ | UKIDSS $J$ | UKIDSS $H$ | UKIDSS $K$ |
|-----------------------|--------|----------|----------|------------|------------|------------|------------|
| ULAS J135122.15+141914.9 | esdM7  | 19.24±0.03 | 18.92±0.03 | 17.01±0.01 | 16.31±0.01 | 15.92±0.01 | 15.63±0.02 |
| ULAS J002009.35+160451.2 | sdM9   | 20.81±0.06 | 19.21±0.05 | 18.07±0.03 | 17.32±0.02 | 16.90±0.04 | 16.65±0.04 |
| ULAS J100756.85+100811.3 | sdM7   | 21.29±0.09 | 20.00±0.11 | 18.80±0.06 | 18.08±0.07 | 17.76±0.08 | 17.50±0.10 |
| ULAS J202648.22+020255.6 | sdM7   | 20.79±0.04 | 19.48±0.05 | 18.45±0.04 | 17.65±0.03 | 17.28±0.07 | 17.10±0.10 |
| ULAS J024035.36+060629.3 | sdM7   | 21.07±0.15 | 19.82±0.08 | 18.66±0.05 | 17.99±0.04 | 17.65±0.06 | 17.48±0.12 |
| ULAS J143154.18−004114.3 | sdM9   | 20.86±0.06 | 19.25±0.05 | 17.92±0.03 | 17.19±0.03 | 16.84±0.04 | 16.48±0.05 |
| ULAS J143517.18−04173.1 | sdM8   | 19.79±0.04 | 18.31±0.03 | 17.20±0.02 | 16.39±0.01 | 15.97±0.02 | 15.62±0.02 |

**Table 6.** Photometric properties of new M–L dwarfs/subdwarfs identified in this work.
| Name            | PS1 name         | $i_{P1}$ | $z_{P1}$ | $y_{P1}$ | $W1$  | $W2$  |
|-----------------|------------------|----------|----------|----------|-------|-------|
| SSSS J013074.34—135620.4 | J01307.404—135622.796 | 17.23±0.01 | 16.25±0.01 | 15.92±0.01 | 13.80±0.03 | 13.60±0.03 |
| J010430.65—143459.325 | J010438.655—143459.325 | 20.52±0.02 | 19.49±0.02 | 19.09±0.03 | 16.61±0.08 | 16.38±0.25 |
| J122537.13—022542.4 | J122537.15—022542.117 | 19.45±0.02 | 18.10±0.02 | 17.51±0.01 | 15.21±0.04 | 15.01±0.08 |
| J130518.85—081506.8 | J130518.782—081506.071 | 21.26±0.17 | 19.76±0.02 | 19.32±0.04 | 17.45±0.20 | >16.45 |
| J162618.76—392519.0 | J162618.820—392522.418 | 17.92±0.01 | 16.34±0.01 | 15.79±0.01 | 13.48±0.02 | 13.14±0.03 |
| WISEA J121409.1—713263.2 | J121409.144—713263.852 | 19.09±0.01 | 17.99±0.02 | 17.58±0.02 | 15.23±0.03 | 14.95±0.04 |
| J230711.01—014447.1 | J230711.008—014446.289 | 21.69±0.16 | 20.16±0.04 | 19.57±0.08 | 17.47±0.18 | >16.79 |
| WISEA J010517.0—083823.4 | J010517.15—083823.247 | 17.40±0.06 | 16.26±0.01 | 15.77±0.01 | 13.43±0.03 | 13.20±0.03 |
| J020201.25—313645.2 | J020201.198—313649.894 | 18.48±0.01 | 17.08±0.02 | 16.57±0.01 | 15.13±0.03 | 14.96±0.04 |
| J020858.62—020657.0 | J020858.625—020656.920 | 21.39±0.03 | 19.99±0.03 | 19.50±0.06 | 17.17±0.13 | 16.63±0.29 |
| J033511.10—00405.4 | J033511.332—00405.465 | 19.22±0.01 | 18.05±0.01 | 17.54±0.01 | 15.08±0.04 | 14.77±0.07 |
| J030601.66—030359.0 | J030601.672—030359.676 | 18.02±0.02 | 16.57±0.01 | 16.09±0.01 | 14.00±0.03 | 13.67±0.04 |
| J043535.68—211508.2 | J043535.926—211507.487 | 18.60±0.01 | 17.15±0.01 | 16.69±0.01 | 14.03±0.03 | 13.81±0.04 |
| J053523.46—824645.6 | J053509.378—824662.210 | 20.31±0.06 | 18.09±0.01 | 16.95±0.01 | 13.82±0.03 | 13.26±0.03 |
| J060616.40—604719.4 | — | — | — | — | — | — |
| ULAS J111429.34—072809.3 | J111429.525—072809.033 | 20.69±0.02 | 19.55±0.02 | 19.06±0.03 | 16.84±0.12 | 16.85±0.39 |
| SDSS J124410.11—236258.5 | J124410.663—236258.064 | 20.40±0.02 | 19.32±0.01 | 18.81±0.02 | 16.68±0.09 | 16.94±0.40 |
| J043535.68—211508.2 | J043535.926—211507.487 | 18.60±0.01 | 17.15±0.01 | 16.69±0.01 | 14.00±0.03 | 13.67±0.04 |
| J053523.46—824645.6 | J053509.378—824662.210 | 20.31±0.06 | 18.09±0.01 | 16.95±0.01 | 13.82±0.03 | 13.26±0.03 |
| J060616.40—604719.4 | — | — | — | — | — | — |
| ULAS J111429.34—072809.3 | J111429.525—072809.033 | 20.69±0.02 | 19.55±0.02 | 19.06±0.03 | 16.84±0.12 | 16.85±0.39 |
| SDSS J124410.11—236258.5 | J124410.663—236258.064 | 20.40±0.02 | 19.32±0.01 | 18.81±0.02 | 16.68±0.09 | 16.94±0.40 |
| J043535.68—211508.2 | J043535.926—211507.487 | 18.60±0.01 | 17.15±0.01 | 16.69±0.01 | 14.00±0.03 | 13.67±0.04 |
| J053523.46—824645.6 | J053509.378—824662.210 | 20.31±0.06 | 18.09±0.01 | 16.95±0.01 | 13.82±0.03 | 13.26±0.03 |
| J060616.40—604719.4 | — | — | — | — | — | — |
| ULAS J111429.34—072809.3 | J111429.525—072809.033 | 20.69±0.02 | 19.55±0.02 | 19.06±0.03 | 16.84±0.12 | 16.85±0.39 |
| SDSS J124410.11—236258.5 | J124410.663—236258.064 | 20.40±0.02 | 19.32±0.01 | 18.81±0.02 | 16.68±0.09 | 16.94±0.40 |
| J043535.68—211508.2 | J043535.926—211507.487 | 18.60±0.01 | 17.15±0.01 | 16.69±0.01 | 14.00±0.03 | 13.67±0.04 |
| J053523.46—824645.6 | J053509.378—824662.210 | 20.31±0.06 | 18.09±0.01 | 16.95±0.01 | 13.82±0.03 | 13.26±0.03 |
| J060616.40—604719.4 | — | — | — | — | — | — |

Table 7. PSI and WISE photometry of L dwarfs.

New L subdwarfs and population properties
objects that were not observed by UKIDSS. We also gathered W1, W2, and W3 photometry from the all-sky WISE data release, and $i_{P1}, z_{P1},$ and $y_{P1}$ photometry from Pan-STARRS (PS1; Chambers et al. 2016). This compilation of optical to infrared photometry of 66 L subdwarfs is presented in Tables 1 and 7.

Some known L subdwarfs that were discovered by 2MASS and WISE were not observed by UKIDSS or VISTA’s Hemisphere Survey (VHS) and Kilo-Degree Infrared Galaxy Survey (VIKING). To convert 2MASS photometry ($J_2, H_2,$ and $K_2$) into the MKO system, we determined polynomial relationships between MKO and 2MASS photometric differences and spectral type (SpT) for known L0–T0 dwarfs. Ideally, these relations would be based on L subdwarf measurements, but data availability limits us to take the next best approach, and use L dwarfs:

$$J - J_2 = -0.0230 - 0.0023 \times \text{SpT}; \quad (0.1062), \quad (1)$$

$$H - H_2 = 0.0188 + 0.0029 \times \text{SpT}; \quad (0.1027), \quad (2)$$

$$K - K_2 = -0.0977 + 0.0068 \times \text{SpT}; \quad (0.0938). \quad (3)$$

where SpT is 10 for L0, 15 for L5, and 20 for T0. The root mean square (rms) of these fits are given after each equation.

Some L subdwarfs were not in the SDSS footprint but were observed by PS1. We thus studied the correlation between SDSS and PS1 photometry. We found that SDSS and PS1 photometry are similar in the $i$-band ($< 0.04$ for main-sequence stars), but correlate with spectral type in the $z$-band. This is presumably because the SDSS and PS1 $i$-band filters have similar wavelength coverage and transmission profiles, though their $z$-band filters are different. These differences between SDSS and PS1 $z$-band magnitudes ($z - 2z_{P1}$) are around $-0.05$ for FGK dwarfs, $-0.11$ for M0–M3 dwarfs, and $-0.43$ for L dwarfs. We determined polynomial relationships between SDSS–PS1 photometric differences and spectral type for known L0–T0 dwarfs:

$$i - i_{P1} = -0.0565 + 0.0043 \times \text{SpT}; \quad (0.0944), \quad (4)$$

$$z - 2z_{P1} = -0.3092 - 0.0086 \times \text{SpT}; \quad (0.0886). \quad (5)$$

Fig. 15 shows our relationships between SDSS–PS1 and MKO–2MASS photometric differences and spectral types for L0–T0 dwarfs.

Four of 66 known L subdwarfs (2M0041, WISEA J005757.65+201304.0, UL1241-00, and WI1355) are excluded from our sample due to problematic photometry or ambiguous spectral type. We thus have 62 L subdwarfs in our photometric sample.

### 4.2 L subdwarf colour–spectral type relations

To identify the best colours for distinguishing and characterizing L subdwarfs with different spectral types and sub-
classes, we assessed various optical to near- and mid-infrared colours. Fig. 16 shows optical to infrared colours of sdL, esdL, and usdL subdwarfs compared to third-order polynomial fits of spectral type – colour relations for L dwarfs. The coefficients of these polynomial fits and rms are presented in Table 8. L dwarfs used for these polynomial fits are collected from DwarfArchives.org, and cross-matched within the SDSS, UKIDSS, PS1, and WISE data bases. Note that a few objects near the WISE detection limit have photometric uncertainties of 0.6 mag.

The \( J - H \), \( J - K \), \( J - W2 \), \( yP1 - H \), \( zP1 - H \), \( yP1 - K \), \( Y - W1 \), and \( J - W1 \) are the best metallicity indicators, as they can be used to separate sdL, esdL, and usdL subclasses from L dwarfs. The \( i - J \), \( z - J \), \( iP1 - yP1 \), \( zP1 - yP1 \), \( iP1 - J \), \( yP1 - J \), and \( iP1 - H \) can be used to separate esdL and usdL from L dwarfs, but cannot distinguish sdL from L dwarfs well.

The esdL subclass has bluer \( H - W1 \) and \( H - W2 \) colours compared to L dwarfs, but the usdL subclass tends to have relatively redder \( H - W1 \) and \( H - W2 \) colours. The most metal-poor usdL subdwarf, SD0104+15 (Paper II), shows a redder \( H - W2 \) and much redder \( H - W1 \) compared to L dwarfs. The esdL and usdL subclasses are redder in \( K - W1 \) than the L dwarfs. Note that SD0104+15 is much redder in \( K - W1 \) than the L dwarfs.

The \( i - z \), \( iP1 - zP1 \), \( H - W2 \), \( W1 - W2 \), and \( W2 - W3 \) colours are not good metallicity indicators. The spectral subtypes of L subdwarfs are based on red optical spectra (\( i \) and \( z \) bands) by comparison to L dwarfs (Paper I). Therefore, L subdwarfs have similar \( i - z \) colours to L dwarfs with the same subtype. However, early-type esdL and usdL subdwarfs have a different optical spectral profile to early-type L dwarfs, and have bluer \( i - z \) and \( iP1 - zP1 \) colours compared to L dwarfs. L subdwarfs also appear to be slightly redder in \( W2 - W3 \) than the L dwarfs.

### 4.3 L subdwarfs in colour–colour plots

Colour–colour plots are often used to distinguish celestial objects of different types in large-scale photometric surveys. Particular objects occupy a variety of colour spaces, with varying level of contamination from other populations (e.g., fig. 2 of Zhang et al. 2013).

Fig. 17 shows the \( i - J \) versus \( J - K \) colour–colour plot for sdL, esdL, and usdL subdwarfs compared to L dwarfs and main-sequence stars. The figure also shows the BT-Dusty model predicted colours of L subdwarfs (Allard, Homeier, & Freytag 2014). The different metallicity ranges of these four subclasses are well represented by their broad-band optical-NIR colours, and the dL, sdL, esdL, and usdL sequences are quite distinct from late M through the L types. The \( i - J \) and \( J - K \) colours are very useful for selecting L subdwarf candidates. The \( i - J \) colour is a good indicator of \( T_{\text{eff}} \) and can be used to separate L subdwarfs from main-sequence stars, while the \( J - K \) colour is a good indicator of metallicity and can be used to separate L subdwarfs from L dwarfs.

Fig. 18 (left) shows the \( z - J \) versus \( J - H \) colour–colour plot for L subdwarfs and dwarfs. The main-sequence population and BT-Dusty model grid with \( g = 5.5 \) are plotted for comparison. The dL, sdL, esdL, and usdL sequences are revealed in the \( z - J \) versus \( J - H \) plane, although not as well as in Fig. 17. Fig. 18 (right) shows the \( J - W2 \) versus \( J - K \) colour–colour plot showing dL, dT, sdL, esdL and usdL populations. The L subdwarf populations are generally separated from each other by their \( J - W2 \) and \( J - K \) colours, although they are very close to the main-sequence. Late-type L subdwarfs begin to merge with the T dwarfs on the right side of the diagram.

To further explore the potential of future L subdwarf searches and characterization, Fig. 19 compares our L subdwarf photometric sample to L dwarfs and main-sequence stars in twelve optical/NIR, NIR, and NIR/mid-IR colour–colour diagrams (constructed using the PS1, MKO, and WISE photometric systems). \( J - K \) or \( J - H \) colour are used for 1 of these plots, and are sensitive to metallicity. As for Figs 17 and 18, subdwarfs of different subclasses in these eleven plots are generally well separated. The longer base-line optical/NIR colours (e.g., \( iP1 - J \) and \( zP1 - J \)) are relatively better at separating objects with different spectral type. The last three colour–colour plots in Fig. 19 use the \( K - W1 \) colour and are very good at isolating the most metal-poor L subdwarfs (usdLs), which have red \( K - W1 \) colour (also see Fig. 16).
Figure 16. Optical to infrared colours of L subdwarfs. Red hexagons, blue circles, and black diamonds represent sdL, esdL, and usdL subclasses. Third-order polynomial fits (black lines) and their rms (green shaded areas) for colour–spectral type relations of L dwarfs are also plotted for comparison.
Figure 17. The $i - J$ versus $J - K$ colours of L subdwarfs. Green open circles, red hexagons, blue circles, and black diamonds represent dL, sdL, esdL, and usdL populations, respectively. Error bars are smaller than the symbol size for some objects. Grey dots are 5000 point sources selected from a 10 deg$^2$ area of ULAS-SDSS-PS1-WISE sky with $14 < J < 16$, which represents main-sequence stars (mostly with spectral types of FGK and M0–M4). The broken dashed line indicates an empirical stellar–substellar boundary (Paper II). The $i$-band magnitudes of UL2307 and VLMS (magenta filled five-pointed stars) in the local field (Dieterich et al. 2014) have been converted from $ip_1$ with equation (4). We converted 2MASS magnitudes to MKO magnitudes for L subdwarfs not observed in UKIDSS. The blue open circle on the right indicates 2M0616 with an estimated $i - J$ colour. The blue open circle in the middle indicates UL0208 with uncertain $K$ band photometry. The BT-Dusty model grid with log $g = 5.5$ is plotted for comparison (with $T_{\text{eff}}$ and [Fe/H] indicated). M and L subtypes are marked at corresponding locations by their average colours. UL2332+12 (magenta open square) and SD1331 (magenta open pentagon) are joined by a magenta dashed line (see Section 3.4).

Figure 18. The $z - J$ versus $J - H$ (left) and $J - W2$ versus $J - K$ (right) colour–colour plots for the L subdwarfs. Symbols are as described in Fig. 17, with additional magenta open squares to represent T dwarfs.
Figure 19. Optical to infrared colour–colour plots for L subdwarfs. Symbols are as described in Fig. 17.
cores which partially contributes to their luminosity (Paper II & III). As the initial thermal energy of a transitional BD is slowly dissipating over time, the unsteady hydrogen fusion slowly becomes the dominate energy source to maintain its luminosity. Since the core temperature of a transitional BD declines slowly over time. The efficiency of the fusion also declines very slowly over time. The fusion is very sensitive to the mass of transitional BDs. Therefore, field transitional BDs at a certain age could span in a wide $T_{\text{eff}}$ range within a narrow mass range. Note such a $T_{\text{eff}}$ range is even wider at older age or lower metallicity. The long-lasting unsteady hydrogen fusion in the cores of transitional BDs slowed down their cooling speed and also made them different from both VLMS and degenerate BDs.

Most of field degenerate BDs would have evolved into the mid- to late-type T dwarf domain by their ages of ~0.5–5 Gyr. The rest of them with relatively younger age or higher mass would be crossing the L dwarf domain. This is the main reason for that L dwarfs have a much lower number density than T dwarfs in the solar neighbourhood (Burgasser 2007; Kirkpatrick et al. 2012). Field L dwarfs are composed of VLMS, transitional BDs, and degenerate BDs which are difficult to distinguish by observation without knowing their mass or age. Field transitional BDs are mostly L dwarfs, but could have spectral types of late-type M and early-type T, depending on their mass and age. The BD transition-zone, which has not drawn much attention in the past, further blurred the observational stellar/substellar boundary in addition to the mass/age degeneracy. This is why we have not reached an agreement on the observational stellar/substellar boundary among field population.

4.5 The observational stellar/substellar boundary

Transitional BDs in the field are difficult to identify by observation because of the mass/age degeneracy and relatively small luminosity/temperature distinction (e.g. GD 165B; Kirkpatrick et al. 1999b). However, transitional BDs in the halo are distributed in a very wide range of luminosity/temperature after over ~10 Gyr of cooling to reduce their initial thermal energy. Nine known halo transitional BDs identified by their atmospheric parameters are summarised in Paper III. They are located in a very broad halo BD transition-zone in the $T_{\text{eff}}$ versus [Fe/H] parameter space.

To identify an empirical observational boundary between metal-poor stars and transitional BDs, we compared known metal-poor transitional BDs to least-massive stars in the halo, thick disc (in fig. 9 of Paper II), and the solar neighbourhood (Dieterich et al. 2014) by their relative locations on colour–colour plots presented in Section 4.3. Fig. 17 shows that metal-poor transitional BDs can be well separated from stars by their $i - J$ and $J - K$ colours that are sensitive to subtype (i.e. $T_{\text{eff}}$) and subclass (i.e. [Fe/H]), respectively. Therefore, we draw an empirical stellar–substellar boundary in Fig. 17 that is shown as a black dashed broken line.

The stellar–substellar boundary was drawn where we can separate known transitional BDs from least-massive stars as much as possible by their $i - J$ and $J - K$ colours. Then we lead this boundary to the right side of the two least-massive field stars suggested by Dieterich et al. (2014).
Figure 21. The Hertzsprung-Russell diagram for L subdwarfs in comparison to field objects. Symbols are as described in Fig. 17. Transitional BDs are indicated with magenta circles. Grey dots are objects selected from Gaia DR2, PS1, and LAS with distance < 100 pc, $180^\circ < RA < 220^\circ$ and $0^\circ < Dec. < 20^\circ$. The two grey sequences are white dwarfs (left) and main-sequence stars (right). Some field stars are scattered mostly because they are too bright in the PS1 and UKIDSS fields.

Eight out of these nine known halo transitional BDs in Paper III lie below/redward of this stellar–substellar boundary. However, we note one apparent inconsistency, since UL0208 appears just above the boundary (blue open circle in the middle of Fig. 17). We believe this is likely due to the uncertain $K$ band photometry of this object (giving a larger $J - K$ error), and further note that this object lies below the boundary in Fig. 18 as a result of its NIR $J - H$ colour. It can also be seen that there are a significant number of esdL stars clustered just above/blueward of the boundary, and considerably fewer esdL subdwarfs below/redward of it. This dearth of objects in the colour space on the substellar side of the boundary represents a substellar subdwarf gap spanning mid L to early T types. This region (the transition-zone) covers a wide $T_{\text{eff}}$ range but a narrow mass range, which results in relatively fewer objects compared to halo VLMS.

We also draw stellar/substellar boundaries on other colour–colour plots composed of $T_{\text{eff}}$ and [Fe/H] sensitive colours (Figs 18 and 19). The corresponding kink points on these boundaries in different plots have the same value for each colour. So that the boundaries in different colour–colour plots are consistent. Metal-poor transitional BDs are well separated from stars across the sdL, esdL, and usdL populations in Fig. 18 and these first nine plots of Fig. 19. The last three plots in Fig. 19 use the $K - W1$ colour could only separate transitional BDs of esdL/usdL subclasses and invalid for sdL subclass.

Using the stellar/substellar boundaries from these plots in Figs 17–19, we estimate that about 22 of these 62 L subdwarfs in our photometric sample are likely metal-poor transitional BDs which are in the substellar subdwarf gap between VLMS and degenerate BDs. Halo degenerate BDs would have temperature below $\sim 1000$ K and have cooled into the T- and Y-type region (Paper II; Burrows et al. 2001). Field degenerate BDs could have much earlier spectral types depending on their ages. Degenerate BDs do not have unsteady hydrogen fusion, but the most massive ones above $\sim 0.055 M_\odot$ (Fig. 20) would be able to fuse lithium. Therefore, ultra-cool dwarfs with lithium absorption lines in their spectra would fall in the degenerate BD domain (e.g. Teide 1; Rebolo et al. 1996). However, lithium absorption lines are not expected in the spectrum of halo degenerate BDs because of their low temperature (section 3.5 of Paper III).

The least-massive field stars above the HBMM have spectral types of around L2.5 (2MASSI J0523382−140302; Dieterich et al. 2014) and L3 (2MASSI J1017075+130839B; Dupuy & Liu 2017). The latest spectral types of stars in our L subdwarf sample are sdL1/esdL1/usdL0. Note we do not have sdL2/esdL2/usdL1 subdwarfs in our sample; thus we do not know if they could be stars, as they would be very close to the stellar boundary if they exist. The subtypes of L subdwarfs are assigned by comparing their optical spectra with those of L dwarf standards. L subdwarfs are hotter, more luminous, and more massive than L dwarfs with the same spectral type. Meanwhile, the HBMM is higher at lower metallicity (Paper II and III). As a consequence of these two facts, the latest spectral types of stars is slightly shifting to earlier subtypes across the dL, sdL, esdL, and usdL subclasses with decreasing metallicity.
Figure 22. Correlations between spectral types and absolute magnitudes. The blue, red, and green lines indicate our polynomial fits for esdL/usdL, sdL, and dL subclasses. Their rms are indicated by shaded areas. The polynomial fits for field dwarfs in J, H and K bands are from Dupuy & Liu (2012). Other symbols are as described in Fig. 17.

5 GAIA OBSERVATIONS OF L SUBDWARFS

The ESA’s Gaia (Gaia Collaboration et al. 2016) astrometric survey have measured precise parallaxes and proper motions (Lindegren et al. 2018) for ~1.332 billion stars in its second data release (DR2; Gaia Collaboration et al. 2018). The Gaia survey has three optical pass-bands ($G$, $G_{\text{BP}}$ and $G_{\text{RP}}$) that are not very sensitive for ultra-cool dwarfs. Ultra-cool subdwarfs is slightly easier to be detected by Gaia than dwarfs with the same spectral types. As ultra-cool subdwarfs have brighter magnitude in the optical (see fig. 3 in Zhang et al. 2013).

Twenty of 66 known L subdwarfs in Table 1 were observed by Gaia. Table 9 shows the Gaia astrometry and photometry of these 20 L subdwarfs. As our UKIDSS-SDSS survey is relatively deep, only two of our L subdwarfs are in Gaia DR2; UL0753+20 is from this work and ULAS J134749.79+333601.7 was presented in Paper I. The rest are previous known L subdwarfs.

To better understand the properties of L subdwarfs and estimate the distance of L subdwarfs that are not in Gaia DR2, we studied the correlation between spectral type and absolute magnitude of L0–7 subdwarfs. Fig. 22 shows the first-order polynomial fits of relationships between their spectral types and Gaia, PS1, and MKO absolute magnitudes.
tudes. Table 10 shows the coefficients of these fittings (see Table 8 for L0–L9 dwarfs). The $G_{BP}$ band magnitudes of known L subdwarfs/dwarfs are likely close to Gaia’s detection limit; therefore, these $M_{G_{BP}}$ in Fig. 22 (b) might not be reliable. In general, the $G$ to $J$ absolute magnitudes get brighter from dL to sdL, and esdM/USDL subclasses. The esdM/USDL subclasses have similar $M_{J}$ to L dwarfs but fainter $M_K$. These are similar to what were shown in fig. 3 of Zhang et al. (2013). However, the $M_K$ of sdL subclass seems fainter at L0 but brighter at L7 than dL, esdM/USDL subclasses. The $M_K$ of sdL subclass is also fainter at L0 but brighter at L7 than esdL/USDL subclasses. Fig. 23 shows the tangential velocities of 20 L subdwarfs observed by Gaia. These field stars in Fig. 23 have an median $V_{tan}$ of ~36 km $s^{-1}$. Five of these seven sdL subdwarfs have $V_{tan}$ > 75 km $s^{-1}$. All the esdL/USDL subdwarfs have $V_{tan}$ > 120 km $s^{-1}$, and most of them are between 200 and 400 km $s^{-1}$. The esdL and USDL subclasses generally have halo kinematics, which is consistent to the esdM/USD subclasses (on the classification of Lépine, Rich, & Shara 2007). The sdL subclass mostly have thick disc kinematics. There are a few relatively more metal-poor ([Fe/H] $\approx$ –1)
sdL subdwarfs have halo kinematics: UL0212+06, SD1333 (Paper I) and VVV J12564163–6202039 (Smith et al. 2018).

6 CONCLUSIONS AND FUTURE DIRECTIONS

We present the discovery of 27 L subdwarfs including 5 esdL and 22 sdL subdwarfs. These new objects were classified according to the L subdwarf classification scheme presented in Paper I. Six of these L subdwarfs have spectral types between L3 and L8 and are likely substellar objects, while the other 21 L0–L1 subdwarfs are likely VLMS. We measured their proper motions and estimated their spectroscopic distances. We also measured the RV of three that have X-shooter spectroscopy. Our SDSS-UKIDSS programme has confirmed/classified 35 L subdwarfs in total (amongst a full known population of 66), including 11 probable BDs and 24 VLMS.

We also interpret one of our candidates (UL2332+12) as a mildly metal-poor unresolved binary consisting of a blue ∼L6p primary and a ∼T4p secondary. UL2332+12 has a high probability of thick disc membership by its kinematics. Metal-poor BD binaries are rare, but their properties and binary fraction may be very useful for our understanding of substellar formation in the early Galaxy (e.g. Bate 2014; Stamatellos & Whitworth 2009).

We have assessed optical to mid-infrared colours of the L subdwarf population, using colour-spectral type and 2-colour diagrams, comparing with both L dwarfs and main-sequence stars. We found that L subdwarfs of different metallicity subclasses can be well separated from L dwarfs and main-sequence stars using a range of optical/infrared colours. Colour spaces have been identified in which preferential selection can be made of the full range of L subdwarf subclasses, as well as separating stellar and substellar subdwarfs. This analysis shows that the photometric systems employed by the PS1, VST and VISTA surveys provide strong future potential for expansion of the known L subdwarf population out to its metallicity extremes, for which (based on the current sample) around a third will be substellar objects with the remaining two-thirds VLMS.

We also note that the PS1 \( p_1 \), \( z_1 \), and \( g_1 \) filters have similar wavelength coverage and transmission profiles to those planned for the next decade’s Large Synoptic Survey Telescope (LSST; LSST Science Collaboration et al. 2017) and the Chinese Space Station Optical Survey (CSS-Os). The LSST will observe the southern hemisphere in six passbands to single-visit depths of 23.4, 22.2, and 21.6 mag and co-added depths of 26.4, 25.2, and 24.4 mag in the \( i \), \( z \), and \( y \) bands, respectively, from 2022. A large area in the northern sky missed by the LSST will be covered by the CSS-Os. The CSS-Os will observe ∼ 17 500 deg² (\( |b| > 15 \) deg and \( |Dec.| > 20 \) deg) of the sky in seven bands to depths of 25.9, 25.2, and 24.4 mags in \( i \), \( z \), and \( y \) respectively. The LSST could provide useful parallax distances for objects well brighter than its single-visit depth using the Gaia’s reference frame. Meanwhile, the CSS-Os have much deeper single-visit depth than the LSST, therefore, is better in detecting higher proper motion cool objects beyond LSST’s single-visit depth.

Furthermore, ESA’s Euclid (Laureijs et al. 2011) space survey telescope is scheduled to launch in 2021, and aiming to observe half of the sky in four pass-bands to depths of 25 mag in VIS band and 24 mag in \( Y \), \( J \) and \( H \) bands. Its slitless spectroscopy will observe the 0.92–1.85 \( \mu m \) wavelength to a depth of \( H \approx 19.5 \) mag and could be used to identify T subdwarfs by their \( Y/J \) index (Burgasser, Burrows, & Kirkpatrick 2006; Mace et al. 2013a). The NASA’s Wide-Field InfraRed Survey Telescope (WFIRST; Spergel et al. 2015) is planning to observe 2000 deg² of the sky in its high-latitude survey to depths of \( Y = 26.7 \), \( J = 26.9 \), \( H = 26.7 \), and \( \text{F184} = 26.2 \) from mid-2020s.

These future optical and NIR sky surveys will provide great opportunity to the study large numbers of extremely metal-poor L subdwarfs and halo degenerate BDs (esdT/Y and usdT/Y types) in the near future. In particular, the \( z \) to \( H \) bands (covered by these facilities) will probe very large volumes for L subdwarfs, with the most extreme examples remaining reasonably bright in these bands.

The \( K \) band flux of ultra-cool objects are the most sensitive wavelength to metallicity disparity. The \( J \)–\( K \) colour is very useful in the identification and characterization of ultra-cool subdwarfs, particularly for T subdwarfs that have similar \( J \)- and \( H \)-band spectral profile to T dwarfs but stronger suppressed flux in \( K \) band (e.g. Burningham et al. 2014). However, the \( K \) band filter is not included in current survey strategies of both Euclid and WFIRST. The Euclid and WFIRST surveys will gain a lot more impacts on the science of ultra-cool objects, if it could include a \( K \) band filter and extend the red cut-off wavelength of their NIR spectroscopy from \( \sim 1.9 \) to 2.2 \( \mu m \). The WFIRST has a larger aperture size than the Euclid, and would have a better capability in the \( K \) band detection of T subdwarfs, which become very faint.

We further discussed the BD transition-zone and properties of transitional BDs following Papers II and III. Degenerate BDs have an essentially different evolution from VLMS but their observational distinction are blurred by transitional BDs. Firstly, because the existence of the BD transition-zone was not widely realised. Secondly, least-massive stars, transitional and degenerate BDs are mixed in the L dwarf domain. Although, in the L dwarf domain, young or massive degenerate BDs are crossing, field transitional BDs are making a long stay or slowly crossing, meanwhile older least-massive stars are more permanent (at early-type L). The BD transition-zone worth in-depth studies by modelling and observation, and that would help us to better understand observations of substellar populations. For example, there is a lack of objects at the L/T transition (Burgasser 2007). This is firstly because the rapid evolution of BD atmospheres at \( \sim 1200 \) K stretched the spectral subtype sampling. Secondly, the L/T transition is at the bottom of the BD transition-zone and next to the abundant degenerate BDs that crossed the BD transition-zone.

Twenty of these 66 known L subdwarfs were observed by the Gaia with precise astrometry. L subdwarfs do appear as ‘sub’ dwarfs on the HRD with some specific colours (e.g. \( z_1 - p_1, z_1 - J \) and \( J - K \)). Their absolute magnitudes are brighter in optical to \( J \) band and fainter in \( K \) band than L dwarfs. The esdL and usdL subclasses generally have halo kinematics and the sdL subclass mostly have thick disc kinematics, which is consistent to M subdwarfs. Five of these 20 L subdwarf in Gaia DR2 are transitional BDs. Degenerate
BDs or T subdwarfs are too faint for the Gaia survey. However, we could get the precise astrometry of a T subdwarf if it has a bright companion observed by the Gaia. For example, WISE J200520.38+542433.9 (Mace et al. 2013b) is an sdT8 subdwarf companion (separated by 188.5 arcsec) to Wolf 1130 that has a Gaia DR2 distance of 16.5587±0.0094 pc. Note that the uncertainty of the Gaia distance of Wolf 1130 is about two-thirds of the projected separation from its cool companion.

More wide binary systems contain both metal-poor BDs and FGKM stars are expected to be discovered in the near future (e.g. Marocco et al. 2017). Such systems can be used as benchmarks to characterize metal-poor BDs and test new ultra-cool atmosphere models and substellar evolutionary models. As we could have precise measurements of distances (from Gaia) and abundances (from high resolution spectroscopy) of these bright FGKM primaries, which could be applied to their metal-poor BD companions.

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