**1. Introduction**

The Juno spacecraft has provided an unprecedented glance into Jupiter's atmospheric flows below the cloud level. The high-precision gravity measurements, particularly those of the odd gravitational harmonics repeated in multiple passes (Jess et al., 2018), have presented an opportunity to estimate the depth and structure of Jupiter's zonal jets. It has been found that the zonal jets are deep and penetrate to approximately 3,000 km below the cloud level (Kaspi et al., 2018). Below this depth, the even gravitational harmonics indicate that Jupiter rotates almost like a solid body (Guillot et al., 2018). However, determining the details of the decay profile with depth poses a significant challenge. Remnants of the zonal flows appear even below 4,000 km, and since the estimation of the electrical conductivity in Jupiter at this depth is at least 10 S m⁻¹ (Nellis et al., 1996; Wicht, Gastine, & Duarte, 2019; Wicht, Gastine, Duarte, & Dietrich, 2019), an interaction between the flow and the magnetic field is expected there (Cao & Stevenson, 2017; Duer et al., 2019; Galanti...
et al., 2017; Moore et al., 2019). Understanding the gravity harmonic signature and the flow structure below the cloud level is thus essential in order to build a better picture of Jupiter's atmosphere.

The gravity field of Jupiter, represented by the gravity harmonics, reflects both the internal density structure and the deep zonal flow structure (Hubbard, 1999; Kaspi et al., 2010). The even gravity harmonics are used to constrain the internal density structures of Jupiter and other gas giants (e.g., Heliod et al., 2010; Hubbard et al., 1974, 1975; Nettelmann et al., 2013). Multiple studies have shown that the higher-order (even) gravity harmonics are sensitive to the outer regions of the planet (e.g., Guillot & Gautier, 2007; Nettelmann et al., 2013; Zharkov & Trubitsyn, 1974). Their exact value is defined by the density distribution throughout the planet and the planet's rotation, composition, shape, mass, and radius. Since for a static gas planet, the odd harmonics are identically zero, any gravitational asymmetry between north and south would indicate a dynamical source generating those asymmetries (Kaspi, 2013). Juno measured with high precision the gravity harmonics up to $J_{10}$, including significant odd values. The measured values and error range are $J_1 = (−4.24 \pm 0.91) \times 10^{-8}$, $J_3 = (−6.89 \pm 0.81) \times 10^{-8}$, $J_7 = (12.39 \pm 1.68) \times 10^{-8}$, and $J_9 = (−10.58 \pm 4.35) \times 10^{-8}$ (Jess et al., 2018). The relation between the density anomaly and the flow (thermal wind balance) allows constraining the deep flow structure within the planet (Kaspi, 2013; Kaspi et al., 2010, 2018). Assuming that the cloud-level zonal wind profile is extended towards Jupiter's interior using a scaling factor, one can find many solutions for the deep flow structure that satisfy all four odd gravity harmonics within the uncertainty range. With the currently available data, Jupiter's deep flow cannot be determined uniquely (Kaspi et al., 2018; Kong et al., 2018), and systematic exploration of the range of the deep flow structure is necessary.

Moreover, the meridional profile of the zonal wind is not necessarily constant with depth. The cloud-level wind itself has a measurement error (García-Melendo & Sánchez-Lavega, 2001; Tollefson et al., 2017; Salyk et al., 2006), and as it extends inward, the profile might vary, although any such variation must be accompanied with a meridional temperature gradient as well. Some evidence for such meridional variations come from the Juno microwave radiometer (MWR) measurements showing that the nadir brightness temperature profile (dominated by the ammonia abundance) becomes smoother with depth (Bolton et al., 2017; Li et al., 2017). Although this measurement does not necessarily correlate with temperature, it does coincide, to some degree, with the zonal wind profile at the cloud level (Bolton et al., 2017) and thus might hint to the vertical variation of the zonal wind profile in the upper 300 km.

Previous work on constraining the deep flow structure was done using all four measured gravity harmonics combined (e.g., Kaspi et al., 2018; Kong et al., 2018). However, an important question is how does each gravity harmonic individually constrain the flow strength at different depths. Here, we examine the individual contribution of each odd gravity harmonic, with emphasis on the depth of influence and the relation to the cloud-level zonal wind profile. In order to provide a systematic analysis, we take a hierarchical approach, in which we increase the level of complexity of the variation of the wind structure and in all cases explore what is the range of solutions that match the gravity measurements. We begin with solutions that are identical to the cloud-level profile and allow only for the vertical decay to vary. Then, we relax the constraint on the meridional profile of the zonal wind and allow variations from the measured cloud-level profile along with the varying vertical decay. Finally, we examine random meridional profiles that are not related at all to Jupiter's measured cloud-level profile and explore the possibility that the interior wind structure, which influences the gravity measurements, is completely different from the cloud-level flow. Following this logic, we also search for solutions with smoother wind profiles that resemble the MWR measurements at 300 km (Channel 1) and calculate the vertical profile of such flows that can match also the gravity data.

The paper is organized as follows: In section 2, we introduce the theoretical background for this analysis, connecting the gravity measurements and the wind profile. In section 3, we present the possible solutions for Jupiter's wind profile, the depth sensitivity obtained by excluding a specific harmonic, and the contribution function of each harmonic. In section 4, we discuss the ability to find solutions for the anomalous gravity field of different meridional profiles, and in section 5, we explore depth-dependent meridional structures, inspired by the MWR measurements. We discuss the dynamical implications of the results and conclude in section 6.

## 2. Methodology

The density distribution within Jupiter is reflected in the zonal gravity harmonics ($J_n$), which describe the external gravitational field of the planet in equilibrium (Zharkov & Trubitsyn, 1974). The gravity harmonics...
can be represented by

\[ J_n = -\frac{1}{MR_j} \int_0^1 \rho r^2 P_n(\mu) d\mu dr, \]  

(1)

where \( M \) and \( R_j \) are Jupiter’s mass and equatorial radius, respectively, \( n \) is the harmonic degree \((n = 2, \ldots, N)\), \( \rho \) is the density, \( r \) is the radial coordinate, and \( P_n(\mu) \) is the \( n \)th Legendre polynomial, where \( \mu = \sin \theta \) and \( \theta \) is the latitude (Hubbard, 1984). The density can be decomposed such that \( \rho(r, \theta) = \bar{\rho}(r, \theta) + \rho'(r, \theta) \), where \( \bar{\rho}(r, \theta) \) is the static component that is determined by the planet’s shape and rotation (Hubbard, 2012) and \( \rho'(r, \theta) \) is the dynamical anomaly representing fluid velocities with respect to the solid body rotation (Kaspi et al., 2010). The zonal gravity harmonics that represent only the dynamical part of the flow \( (\Delta J_n) \) can be calculated by integrating the density anomaly and its projection onto the Legendre polynomials in spherical coordinates such that

\[ \Delta J_n = -\frac{2\pi}{MR_j^2} \int_0^1 \rho'(r, \mu) r^{n+2} P_n(\mu) d\mu dr. \]  

(2)

Since an oblate planet with no dynamics is symmetric between north and south, the density anomaly represented by the odd harmonics \((n = 3, 5, \ldots)\) should be identically zero if the flow pattern is symmetric and will be very small if the dynamics are shallow \((\Delta J_n = J_n \text{ for odd } n)\). However, Juno measured four odd gravity harmonics (Iess et al., 2018), indicating the existence of a strong asymmetric pattern in Jupiter’s flow field due to the existence of strong, deep winds.

The rapid rotation and size of the planet (small Rossby number) imply that this asymmetry is directly related to zonal flows, since, to first order, the leading balance in Jupiter is a geostrophic balance between the flow-related Coriolis forces and the pressure gradients. This leads to a vorticity balance known as thermal wind balance (Kaspi et al., 2009; Pedlosky, 1987). If only zonal (azimuthal) flows are considered, the thermal wind balance can be written as

\[ 2\Omega r \frac{\partial \bar{\mu}}{\partial z} = g_0 \frac{\partial \bar{\rho}}{\partial \theta}. \]  

(3)

where \( \Omega \) is Jupiter’s rotation rate, \( \bar{u}(r, \theta) \) is the zonal flow, \( g_0(r) \) is the mean gravitational acceleration, and \( z \) is the direction parallel to the rotation axis. An equivalent equation can be written with temperature instead of density gradients, and one can easily switch between the two versions through the equation of state. Note that the barotropic limit of Equation 3 is not simply when the rhs vanishes, but when \( \frac{\partial u}{\partial z} = 0 \) (see full derivation at Kaspi et al., 2016). Galanti et al. (2017) showed that a higher-order expansion, beyond thermal wind, only slightly adjusts the deep flow dynamics (less than 10%). Therefore, for the purpose of studying the overall vertical profile, Equation 3 is a good approximation.

Our goal here is to search for possible deep wind structures that can explain each of the measured odd gravity harmonics \((J_3, J_5, J_7, \text{ and } J_9)\). Unlike previous studies (e.g., Kaspi et al., 2018), we are not solving for an optimal solution with respect to the full error covariance matrix. Any vertical wind profile that fits the odd measured gravity harmonics, within the uncertainty range of Juno, is considered a possible solution for the flow. This allows us to examine the full range of possible solutions, without converging on a single decay profile of the flow. For example, the optimal solution suggested by Kaspi et al. (2018) that considered the error covariance matrix is not a solution here since the value of \( J_3 \) is not within the measured error.

3. The Vertical Profile of the Zonal Flow

Taking a hierarchical approach entailing an increasing level of complexity, we first use the observed cloud-level wind as an upper boundary condition for the flow field and assume the same profile continues inward in a direction parallel to the spin axis, due to angular momentum considerations (Busse, 1976; Kaspi et al., 2010). The possible deep flow structures are then set to decay continuously from the cloud level to a few thousand kilometers below it (Kaspi et al., 2018), using two different decay regions. Dividing the decay functions into two distinct regions stems from the magnetic field’s possible effects on the flow, expected approximately at \( r < 0.97 R_j \) (Duer et al., 2019; Wicht, Gastine, Duarte, & Dietrich, 2019), which imply that once the electrical conductivity becomes dominant, the magnetic field acts to dissipate the flow (Gastine
et al., 2014; Liu et al., 2008). Thus, for the lower part (the semiconducting region), we chose an exponential decay (Equation 5, \( r < R_T \)) that fits the exponential nature of the electrical conductivity within Jupiter (French et al., 2012; Nellis et al., 1992; Weir et al., 1996). For the upper part, the vertical decay function includes both an exponent and hyperbolic tangent (Equation 5, \( R_T \leq r \leq R_J \)), which combine to give a wide range of possible decay profiles.

The vertical profile of the zonal flow is defined with six independent parameters, chosen to cover an extensive range of vertical profiles. It is set as

\[
\begin{align*}
Q_s(r) &= \begin{cases} 
(1 - \alpha) \exp\left(\frac{r - R_J}{H_1}\right) + \alpha \left[ \frac{\tanh\left(\frac{R_T - r}{H_2}\right)}{\tanh\left(\frac{R_J - r}{H_2}\right)} + 1 \right] & R_T \leq r \leq R_J \\
Q_s(R_T) \exp\left(\frac{R_T - R_J}{H_3}\right) & r < R_T,
\end{cases}
\end{align*}
\]

where \( u(\theta, r) = u_{proj}(\theta, r)Q_s(r) \).

From the 5 \( \times \) 10^5 decay options examined, 6,712 vertical profiles are compatible with Juno’s measured odd gravity harmonics, which represent a little over 1% of the sample population. All the compatible decay profiles are located in a relatively narrow envelope, especially in the region around 2,000 km depth and the one below 4,000 km, with all the options pointing to remnants of jet-associated velocities at a depth
of 4,000 km (Figure 1). Those deep velocities are still on the order of 1 m s\(^{-1}\) and despite being small, they are still higher than the magnetic secular variation associated velocities estimated by Moore et al. (2019).

Increasing the error range of Juno’s gravity measurements does allow for more solutions, but the overall structure does not change much (Figure 1b).

### 3.1. The Depth Sensitivity of the Odd Harmonics

Research to date has focused on finding vertical profiles that match all four odd gravity harmonics. However, there is information to be obtained from each gravity harmonic separately. Here, vertical flow profiles that fit three out of the four measured odd gravity harmonics are considered, and the depth sensitivity of the excluded harmonic is studied by examining the difference between the vertical profiles that include the specific \(J_n\) to those that do not necessarily include it. The resulting depth sensitivity of each odd measured gravity harmonic, according to Jupiter’s measured zonal profile, is presented in Figure 2. The gray envelope, the same one as in Figure 1b, is the boundary of all the possible solutions that fit all four odd gravity harmonics within 3\(\sigma\). Note that not all possible profiles inside the gray envelope will necessarily generate a solution compatible with the measured gravity field, since the solution is also dependent on the decay profile within the given envelope. All the other solutions gained while excluding one of the odd gravity harmonics appear in Figure 2 (turquoise envelopes). The turquoise envelopes always contain the gray envelopes by definition, since they are constructed by fitting at least three gravity harmonics. The difference between the turquoise envelopes and the gray ones denote the region in which the excluded harmonic bounds the flow.

The most insignificant influence is clearly of \(J_9\) (Figure 2d). It appears to add no solutions at all to the gray envelope, meaning that \(J_9\) does not constrain the flow if the other three odd values are still within Juno’s 3\(\sigma\). This is likely because \(J_9\) has the highest measurement error and lowest signal-to-noise ratio (SNR), so even while fitting \(J_9\), there is an extensive region of solutions, and excluding it does not add new solutions. The largest influence on the flow profile and depth sensitivity comes from \(J_5\) (Figure 2b). It appears to set the upper boundary of the gray envelope from the cloud level (0 km) to 3,500 km and a lower boundary of the gray envelope between 2,000 and 3,500 km. The strongest sensitivity is between the cloud level and 3,000 km. \(J_5\) has the smallest measured 3\(\sigma\) value and largest SNR, but its value is very similar to the SNR of \(J_7\), so the large influence of \(J_5\) cannot be a result of the SNR alone. In a similar manner, \(J_7\) is mostly sensitive...
Figure 3. (a) The mean anomalous density profile of all possible decay options that fits the Juno four measured odd gravity harmonics; colors represent the anomalous density values (kg m$^{-3}$). (b) Averaged contribution function (lines) (m$^{-1}$) for each of the odd gravity harmonics and their associated standard deviation (shading); triangles represent the depth of the mean anomaly. Both panels are for all the latitudes and for only the upper $\sim 6,000$ km, below which the anomalous density is 0.

between 3,000 and 5,000 km and between the cloud level (0 km) and 1,500 km (Figure 2a). Note that a flow profile that decays to zero at 4,000 km ($\sim 0.94 R_J$) cannot fit $J_3$. $J_7$ is sensitive between 500 and 2,500 km and sets mainly the lower boundary of the gray envelope at those depths (Figure 2c).

Previous studies that examined the depth dependency of the even gravity harmonics (resulting from the shape and density of a solid-body model, without differential flows) concluded that higher-order harmonics are more sensitive to the density in the outer regions (e.g., Guillot & Gautier, 2007; Nettelmann et al., 2013; Zharkov & Trubitsyn, 1974). This is implied by the radial dependence of the gravity harmonics (Equation 1). However, the above analysis shows that the wind-induced odd harmonics’ depth dependency is more complicated. One exception is $J_3$, the only harmonic that is sensitive below 4,000 km, which resembles the even harmonics’ depth tendency, where the low-order harmonics are more sensitive in deeper regions.

3.2. The Contribution Function

The depth sensitivity of the gravity harmonics can also be examined by calculating directly the depth dependence of $J_n$, defined as the contribution function. This function was calculated in past studies for the even harmonics of Jupiter and other planets (e.g., Guillot & Gautier, 2007; Helled et al., 2010; Nettelmann et al., 2013). The contribution of each shell is the normalized integrant of $J_n$, defined as

$$C_n = \frac{1}{J_n} \frac{dJ_n}{dr} = \frac{1}{J_n} \frac{-2\pi}{MR^2} \int_{-1}^{1} \rho(r, \mu) r^{n+2} P_n(\mu) d\mu$$

(Hubbard, 1984; Hubbard et al., 1974; Zharkov & Trubitsyn, 1974). The even harmonics in past studies were calculated from the background density (solid body models), while in our study we use the wind-induced anomalous density field to calculate the odd harmonics’ contribution, taking $\rho'$ instead of $\rho$ in Equation 6.

The averaged anomalous density profile of all possible decay structures, that are consistent with the four odd gravity harmonics, is presented in Figure 3a. The anomalous density reveals a change of sign at 2,000 km. The averaged odd contribution functions ($C_n$) and standard deviations of each odd gravity harmonic (Figure 3b) corresponding to the solution envelope from Figure 1 show a consistent sign change. Note that the change of sign is exhibited only by the anomalous density, corresponding to the wind shear with depth, and, therefore, does not exist when examining the even harmonics resulting from the static density (e.g., Nettelmann et al.,...
2013). The integrals of the nonnormalized contribution curve, $C_n$, are the gravity harmonic values, $J_n$, so the sign and value of $J_n$ are set by the difference between the positive and negative curves (above and below 2,000 km, not shown). For the averaged anomalous density, the gravity harmonics are: $J_3 = -4.29 \times 10^{-8}$, $J_5 = -7.50 \times 10^{-8}$, $J_7 = 10.8 \times 10^{-8}$, and $J_9 = -6.69 \times 10^{-8}$.

The contribution function reveals a complex depth dependence for all four gravity harmonics. The depth sensitivity of each contribution function is marked by the triangles (Figure 3b), which represent the depth of the mean absolute anomaly. The contribution function of $J_3$, $C_3$, has the largest areas under the curves at both the shallower (0–2,000 km) and deeper regions (>2,000 km). The depth of the mean anomaly, which here equals 2,020 km (Figure 3b, blue triangle), is near the depth of the sign change (2,000 km), meaning that $J_3$ gets near-equal anomalies from both regions. The standard deviation of $C_3$ (Figure 3b, blue shading) is the largest, implying a large variability of the solutions with depth when considering the $J_3$ value. The mean anomaly of $C_3$ is located in the deeper part of the domain (Figure 3b, red triangle), and the standard deviation of $C_3$ is substantial only between 2,000 and 4,000 km. $C_3$ is the only harmonic dominated mostly by the deeper region, emphasizing the important effect of $J_3$ on the deep wind structure (section 3.1). The mean anomalies of $C_7$ and $C_9$ are clearly located in the shallower region (yellow and green triangles <2,000 km), and their standard deviation is small everywhere. The contribution of both $C_7$ and $C_9$ is 0 below 3,000 km, corresponding to Figure 2. Since $C_7$ and $C_9$ lay almost on top of each other, $C_7$ might mask the depth dependency of $C_9$, as revealed in Figure 2 (along with the low SNR of $J_9$), so that if $J_9$ is within Juno’s error range, so is $J_9$. It is evident that the contribution function of the odd harmonics exhibits a more complicated pattern than the classical even harmonics (e.g., Guillot & Gautier, 2007; Helled et al., 2010; Nettelmann et al., 2013). As in the previous analysis, we find that the higher-order odd harmonics are not simply more pronounced in the outer regions. The projection of the wind patterns onto different depths is reflected in the odd harmonics’ contribution at those depths, suppressing the $(r/R) n$ dependency, which is the prominent feature of the even harmonics’ contribution.

### 4. Sensitivity to the Meridional Profile of the Zonal Flows

Next, we relax the assumption, used in the previous section, that the meridional profile of Jupiter’s zonal flow remains constant at all depths. First, the zonal wind profile is measured by tracking cloud motion, which itself has some uncertainty (Tollefson et al., 2017). Second, and most importantly, the assumption that the cloud-level profile extends perfectly to depth along the direction of the spin axis requires the flow to be locally close to barotropic (in the upper few thousand kilometers), which is not necessarily the case. Although the flow cannot be completely barotropic if $Q_s ≠ 1$ (Equation 4), the horizontal density gradients required to balance the vertical changes associated with $Q_s$ may be small (Equation 3). Any further deviation from close-to-barotropic flow must be supported by horizontal density (or temperature) gradients, which themselves must be maintained by some internal mechanism (Showman & Kaspi, 2013). Internal convection models support the scenario that there may be internal shear over the upper few thousand kilometers, but the overall structure of the flow does not change much (Jones & Kuzanyan, 2009; Kaspi et al., 2009). Any significant deviation from the zonal wind profile observed at the cloud level requires significant shear and, therefore, notable horizontal thermal gradients (thermal-wind balance). As this is an open question, for the purpose of this analysis, we examine several cases of zonal wind meridional profiles, under the assumption that the wind profile possibly varies close to the cloud level and then projects inward without further modifications. For the purpose of the gravity analysis, this means that the altered meridional profiles occupy enough mass to affect the gravity field, and the flow observed at the cloud level is limited to a shallow-enough layer so it does not affect the gravity field.

The simplest case is clearly to use the measured profile at Jupiter’s cloud level and allow its magnitude to decay with depth (section 3). A slightly less constraining option is to insert a perturbation into the measured profile, thereby keeping the general form and allowing a varying level of modifications to the cloud-level flow. The perturbed winds chosen here might represent the measured uncertainties in Jupiter’s cloud-level wind (García-Melendo & Sánchez-Lavega, 2001; Tollefson et al., 2017). Finally, random meridional profiles of the zonal flow with a spectra generally similar to that of Jupiter are examined as well.

The modified zonal flow profile is chosen at the cloud level and projected inward along the rotation axis ($u_{proj}$ Equation 4) with a range of vertical profiles, as described in section 3. The profiles are calculated by adding sinusoidal perturbations to the measured wind. The modified profiles have a standard deviation
Figure 4. (a) One hundred examples of the 1,000 perturbed wind profiles (colors, \([\text{m s}^{-1}]\)) and Jupiter's measured wind profile (black). (b.1–b.4) The odd gravity harmonics of the perturbed wind profiles depth sensitivity summary as in Figure 2. Each profile is examined with the same set of decay options. The results shown here are for all meridional and vertical options combined.

\[
\epsilon(\theta) = \sum_{n=1}^{10} \left[ a_n \sin (2n\theta) + b_n \cos (2n\theta) \right],
\]  

(7)

where \(\epsilon\) is the perturbation and \(a_n\) and \(b_n\) are random numbers that are normally distributed around 0 with a standard deviation of 2 m s\(^{-1}\). We first examine 1,000 modified profiles, each constructed by adding the perturbation to the measured wind (section 4.1). In addition, 1,000 random profiles are constructed purely from the \(\epsilon\) function (Equation 7), where \(a_n\) and \(b_n\) have a standard deviation of 30 m s\(^{-1}\). These profiles represent internal winds that are completely unrelated to the observed cloud-level winds (section 4.2).

4.1. Perturbed Cloud-Level Wind Profiles

The perturbed wind profiles (Figure 4a, colors) result in a substantially bigger solution envelope (Figures 4b.1–4b.4, gray) than the one from the measured zonal wind profile case, consistent with the fact that a wider range of wind profiles is allowed. Note that the overall shape has changed and that the flow can even vanish at \(\sim 2,500\) km. This might have an important implication, since the initial time-dependent magnetic field results from Juno imply that the wind at these depths are very weak (Duer et al., 2019; Moore et al., 2019). An important result is that even for the perturbed winds there are no solutions that fit at least three odd \(J_n\) that vanish above 2,000 km. The depth sensitivity of each harmonic is less pronounced than for the measured wind case. This reflects the fact that Figure 4 is a combination of all the possible solutions from 1,000 examined meridional structures. Overall, \(J_5\) is still sensitive in the deeper regions (exemplified by the mean profile being weaker at depth, red line Figure 4b.1), although \(J_7\) and \(J_9\) contribute at depth as well. \(J_9\) turns out to be the most insignificant harmonic and \(J_9\) does affect the depth range of 1,500–2,000 km.
Figure 5. Summary of the solutions for the three presented cases of wind structures: Jupiter's measured wind at the cloud level (blue), 1,000 slightly modified meridional structures (red) and 1,000 random meridional profiles with a similar general structure to Jupiter's meridional profile (orange). The ordinate is a logarithmic scale of percentage relative to all the examined cases. The particular requirement of the solution to match the different odd gravity harmonics is presented by the abscissa.

unlike in the unperturbed wind case. The substantially larger range of solutions, however, does not manifest in more solutions relative to the examined cases. From 1,000 examined profiles individually tested with the decay sample population, only about 0.1% fit the anomalous gravity field compared to about 1% in the unperturbed case (Figure 5, red and blue). This suggests that although perturbed cloud-level wind profiles are possible, it is statistically more likely that a profile that is similar to the projection of the observed cloud-level wind is indeed the profile in the deeper atmosphere of Jupiter.

4.2. The Possibility of Other Zonal Wind Profiles

Next, we consider profiles that do not resemble Jupiter's cloud-level winds (Figure 6a). The resulting solution envelopes of the other zonal wind profiles are relatively similar to the previous case of perturbed winds (not shown). Only a very small subset of profiles (13 meridional profiles out of 1,000, about 1%) fit the four measured odd gravity harmonics (Figure 6b.1). The possibility of fitting two or more odd harmonics is rare and exists in only 7% or less of the zonal wind meridional profiles examined (Figures 6b.2 and 6b.3). $J_3$ is the harmonic that is pronounced in the majority of profiles (Figure 6b.4). In 14% of the examined random profiles, no odd harmonic is within the sensitivity range. In general, the measured harmonic’s alignment with the zonal flow structure does not appear to be coincidental. These findings are expected, considering that it is unlikely that an utterly different profile of zonal profiles arise below the cloud level of Jupiter.

The ability of the 1,000 examined random profiles, each with its sample of decay options, to fit all four odd gravity harmonics is considerably smaller than previous cases—only about 0.01% (Figure 5, orange). This indicates that fitting all four odd harmonics is difficult with random meridional profiles of zonal wind. A summary of the examined cases appears in Figure 5. Note that the ordinate is a logarithmic scale and that 100% stands for all the zonal profiles (1,000 zonal wind profiles other than the measured cloud-level wind) and all decay options ($5 \times 10^5$) for each case. We find that the envelope of possible solutions from Figure 1 stands for $\sim 1\%$ of the tested vertical profiles for zonal flows. The fitting percentage decreases with increasing perturbations, and drops rapidly when switching to random profiles. This trend repeats for all variations of at least three odd harmonics. For all cases, the random winds show a significantly lower fitting percentage than the other cases. We further present the fitting percentage obtained following the exclusion of two and three harmonics. In summary, we find that other meridional profiles of the zonal wind are possible, but
Figure 6. (a) Thirty examples of the 1,000 random zonal wind profiles examined (colors, [m s$^{-1}$]) and Jupiter’s measured wind profile (black). (b.1–b.4) Summary of the random meridional profiles’ correspondence to the odd gravity harmonics. Only $\sim$1% of the zonal profiles fit all four odd gravity harmonics (b.1), 5% of the zonal profiles fit at least three odd gravity harmonics (b.2), 28% of the zonal profiles fit at least two odd gravity harmonics (b.3), and 14% do not fit any of the odd gravity harmonics (b.4). The full compatibility distribution is detailed in the figure.

they are statistically unlikely. This result implies that the meridional profile of Jupiter’s zonal winds extends into the interior along the direction of the spin axis and weakens with depth, and is likely not significantly different from the cloud-level profile.

5. Zonal Wind Profiles Inspired by the MWR Measurements

In addition to the gravity measurements, Juno’s six-channel MWR measurements might also reveal information about the structure of the wind below the cloud level. These measurements are used to calculate the nadir brightness temperature ($T_b$), a profile determined by the opacity of the atmosphere, which in Jupiter is determined mostly by ammonia abundance (Li et al., 2017). The MWR measurements reveal considerable variation of $T_b$ with latitude and depth (Bolton et al., 2017) (Figure 7a, black lines). These variations with depth and the potential relation between $T_b$ and the zonal jets (Ingersoll et al., 2017) suggest that Jupiter’s zonal jets might be depth dependent, similarly to $T_b$, instead of simply projected inward (as in sections 3 and 4).

One approach for describing the relation between the brightness temperature and the zonal jets is taking the brightness temperature as simply temperature. Then, the relation is described by the thermal wind balance, as discussed in section 2. This approach, however, results in equatorial wind that is greater by 2 orders of magnitude than the measured cloud-level wind, which is unrealistic (Bolton et al., 2017). Another approach is taking the brightness temperature as an indicator for ammonia concentration (Ingersoll et al., 2017) and examining the relation to the zonal jets. As an example, such a relation is expected in the meridional circulation (Ferrel cells), where the cell-associated vertical velocity redistributes the substance and is accompanied by zonal jets (Fletcher et al., 2020). Here, we take the latter approach, analyzing a range of depth-dependent meridional profiles, compatible with the brightness temperature variations with depth.
Figure 7. (a) Jupiter's projected wind velocities (colors, [m s\(^{-1}\)]) between latitudes \(-50^\circ\) and \(50^\circ\) in the upper 300 km of Jupiter (left ordinate) combined with nadir brightness temperature lines from Juno's PJ1 (black, right ordinate) in Channels 1 to 6, associated with frequencies of 0.6, 1.2, 2.6, 5.2, 10, and 22 GHz, respectively. (b.1–b.6) Jupiter's projected wind velocities (m s\(^{-1}\)) at Channel 1, 240 bar (b.1), Channel 2, 30 bar (b.2), Channel 3, 9 bar (b.3), Channel 4, 3 bar (b.4), Channel 5, 1.5 bar (b.5), and Channel 6, 0.6 bar (b.6) for a running average of \(\Delta \theta = 0^\circ\), \(\Delta \theta = 4^\circ\), and \(\Delta \theta = 8^\circ\) (darker blue with increasing \(\Delta \theta\), left ordinate). Also shown is the brightness temperature [\(^\circ\)K] (black, right ordinate). The radial projection of the winds with no running average (dashed gray) is also presented in (b.1).

When examining the correlation between \(T_b\) and the zonal jets, a different analysis should be taken at different latitudes and perhaps at different depths. If the zonal jets are associated with multiple Ferrel cells in alternating directions associated with regions of momentum convergence (eastward jets) and divergence (westward jets), a correlation is expected between the zonal velocity and the ammonia concentration gradient. In such a scenario, the meridional cells advect the ammonia concentration, maximizing its gradient where the jet peaks (Fletcher et al., 2020). However, momentum fluxes converging at the equator would lead to a superrotating jet (Kaspi et al., 2009) and might also lead to a maximal ammonia concentration. Therefore, at the equator, the zonal velocity is associated with the concentration itself and not with its gradient.

Table 1
Correlation Coefficients Between \(\nabla T_b\) and the Wind Velocity at Each Channel, for Winds With No Running Average at Depth (\(\Delta \theta = 0^\circ\)), and Correlation Coefficients Between \(T_b\) and the Wind Velocity at the Same Depths With No Running Average (\(\Delta \theta = 0^\circ\)), With Running Average of \(4^\circ\) (\(\Delta \theta = 4^\circ\)), and With a Running Average of \(8^\circ\) (\(\Delta \theta = 8^\circ\)).

| Channel | \(u\) versus \(\nabla T_b\) | \(u\) versus \(T_b\) |
|---------|----------------|----------------|
|         | \(\Delta \theta = 0^\circ\) | \(\Delta \theta = 0^\circ\) | \(\Delta \theta = 4^\circ\) | \(\Delta \theta = 8^\circ\) |
| 1       | 0.05 | 0.64 | 0.76 | 0.86 |
| 2       | 0.02 | 0.74 | 0.80 | 0.84 |
| 3       | 0.01 | 0.63 | 0.65 | 0.68 |
| 4       | 0.15 | 0.41 | 0.41 | 0.42 |
| 5       | 0.24 | 0.13 | 0.13 | 0.13 |
| 6       | 0.42 | 0.10 | 0.10 | 0.10 |

Note. For \(u\) versus \(T_b\), the correlation increases with depth (or decreases with channel) and with running average.
Figure 8. (a) The ability of the depth-dependent wind profiles to fit all four odd gravity harmonics (percentage of solutions) as a function of the smoothing factor in degrees (blue line). The three cases shown in Figure 7b are denoted by red dots. (b) The gravity harmonics distribution for the red dots is compatible with the three case studies in Figure 7b. The ordinate is a scale of percentage relative to the $5 \times 10^5$ decay options examined.

These simple considerations motivates us to examine the correlation both between $u$ and $\nabla T_b$ and between $u$ and $T_b$ (Table 1, Columns 2 and 3). Note that, as in all our experiments, the zonal jets are projected inward along the spin axis, as in section 3 (Figure 7a, colors). It is evident that the correlation between $u$ and $T_b$ is weak at the cloud level (Channel 6, 0.6 bar) but becomes stronger with depth (maximum at Channel 1, 240 bar), while the correlation between $u$ and $\nabla T_b$ is strong at the cloud level and weakens with depth (at Channels 1–3, the correlation is weak). This alone might indicate two opposite meridional cells, one stacked on top of the other (Showman & de Pater, 2005; Fletcher et al., 2020). At the cloud level, the correlation between $u$ and $\nabla T_b$ improves if we do not consider the equatorial region, which is consistent with the Ferrel cells hypothesis. Projecting the winds in the radial direction instead of along cylinders does not improve the correlation to neither $T_b$ or $\nabla T_b$ (Figure 7b.1).

The dominant feature leading to the strong correlation between $u$ and $T_b$ at Channel 1 is the equatorial anomaly, ascending at $\sim 15^\circ S$ and descending at $\sim 15^\circ N$ (Figure 7b.1). While at the cloud level, both the zonal jets and $T_b$ reveal alternating patterns (Figure 7b.6), the waviness of $T_b$ vanishes at deeper depths (Figures 7b.1 and 7b.2). Since $T_b$ is depth dependent, getting smoother with depth from Channel 6 (cloud level) to Channel 1 ($\sim 300$ km depth), we examine zonal jets that are depth dependent. Note that winds, projected along the spin axis, maintain their meridional profile with cylinders, and without further assumptions, are not depth dependent.

The modified (smoothed) wind at Channel 1 is composed using a running average of $\Delta \theta$ degrees latitude, where $\Delta \theta = 0, 1, 2, \ldots, 10^\circ$ (0° means that no running average is applied). The wind at Channel 6 is the observed cloud-level profile; between Channel 1 and Channel 6, the wind strength is interpolated. Finally, the wind profile at the depth of 300 km (Channel 1) is projected inward along the spin axis with a decay profile as in the previous sections without further assumptions. In addition to the projected winds with no depth dependency ($\Delta \theta = 0^\circ$, Figure 7b, light blue), we examine the correlation between $u$ and $T_b$ for two
cases of a depth-dependent zonal wind ($\Delta \theta = 4^\circ$ and $\Delta \theta = 8^\circ$, Figure 7b, blue). Increasing the running average at depth improves the correlation at Channels 1–3 (Columns 3–5, Table 1), implying that the latitudinal variability of the jets might weaken beneath the cloud level.

Next, we examine the ability of the depth-dependent zonal profiles to explain the measured odd gravity harmonics. We examine a range of case studies, from slightly to largely modified depth-dependent profiles, until no solutions are found (Figure 8a). For slightly smoother profiles (small $\Delta \theta$), the ability to fit all four odd $J_n$ is similar to that without any smoothing ($\Delta \theta = 0^\circ$) (Figure 8a). Applying additional smoothing (increasing $\Delta \theta$) decreases the ability to fit the four odd $J_n$. Using more than a $10^\circ$ running average results in no solutions for the odd gravity harmonics. The three case studies ($\Delta \theta = 0^\circ$, $\Delta \theta = 4^\circ$, and $\Delta \theta = 8^\circ$) show a consistent trend when excluding one of the odd harmonics, such that the ability to fit the gravity measurements is reduced when it comes to smoother deep profiles (Figure 8b). This result is compatible with the previous case (section 4), indicating that deep wind that resembles the cloud-level wind can fit the gravity data, and changing the zonal wind structure considerably limits the ability to find a solution. The main conclusion of the analysis presented above is that wind profiles correlated with $T_b$ at depth ($\Delta \theta = 4^\circ$) can adequately fit the gravity measurements.

### 6. Discussion and Conclusions

The main challenge of interpreting the Juno gravity measurements is that the measurements provide only a handful of numbers (gravity harmonics), while the meridional and vertical profiles of the interior flow have many degrees of freedom. Therefore, by definition, the problem is ill posed. Acknowledging this inherent issue, Kaspi et al. (2018) used 4 degrees of freedom for the vertical flow profile (matching the number of the four odd harmonics) and found the best optimized profile for this allowed range. They addressed the nonuniqueness by showing the statistical likelihood of wind profiles for the interior that are completely different than the cloud-level flow. Kong et al. (2018) highlighted the nonuniqueness issue by showing that two different flow profiles can still satisfy the gravity measurements. In this study, we take a more methodological approach and consider a wider range of solutions and analyze their statistical likelihood. The flow profiles we consider, both for the meridional and vertical profiles, are bound by physical considerations. We also address two main issues: First, all previous studies looked at all four odd gravity harmonics together and found the flow profiles best matching all four. Here, we investigate how each one of them separately bounds the flow. Second, in an attempt to coincide the gravity and microwave data, we explore whether deep profiles that are smoother than those of the cloud level, as possibility indicated by the Juno microwave measurements, can be consistent with the gravity measurements.

By assuming that the cloud-level wind profile is projected inward parallel to the spin axis, with some decay profile, we identify the envelope of possible solutions (Figure 1). We then relax the dependence on each of the odd gravity harmonics separately and analyze their individual contribution to the vertical profile of the zonal wind (Figure 2). We find that $J_3$, the lowest-order odd harmonic that represents the dynamics of Jupiter, is sensitive at depths where the conductivity rises (beyond $\sim 3,000$ km), and the magnetic field might be interacting with the flow, resulting in the Lorentz force playing a key role in the dynamics. $J_5$ appears to be the most sensitive harmonic, giving a robust constraint on the vertical profile of the zonal flow alone (Figure 2b). Interestingly, $J_6$ does not give any new constraint on the flow if the other three harmonics are within the sensitivity range (Figure 2d). A possible explanation for the unique nature of $J_5$ comes from exploring the contribution function, which revealed that $J_5$ is most sensitive in the deeper regions, below $2,000$ km (Figure 3).

The modified zonal flow analysis revealed a substantially bigger possible solution envelope than that obtained by extending the cloud-level wind (Figure 4). This implies that the zonal wind's structure may influence the depth sensitivity of each harmonic. Even for the perturbed winds, the flow cannot vanish at depths shallower than $2,000$ km. The case with random winds implies that, with high probability, the wind cannot alter completely below the cloud level. Fitting the four odd gravity harmonics (or three if we ignore $J_6$) requires either similar winds to the measured ones at the cloud level, that would penetrate a few thousand kilometers into the planet, or a very specific and statistically unlikely combination of a meridional and a decay profiles (Figures 5 and 6). Finally, the gravity harmonics induced by the slightly modified depth-dependent meridional profiles, which have a better correlation with the MWR measurements at depth...
(Figure 7 and Table 1), are still within Juno's gravity measurements' uncertainty, indicating that Jupiter's ammonia abundance could indeed reflect the profile of the zonal jet at 300 km (Figure 8).

**Data Availability Statement**

The Juno gravity measurements and MWR measurements are publicly available, see https://pds-atmospheres.nmsu.edu/data_and_services/atmospheres_data/JUNO/juno.html. Additional data can be found here https://doi.org/10.5281/zenodo.3859828.

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