### Table 1. Mira Variables

| Star     | RA (2000) | Dec (2000) | Sp. Type | Period (d) | Phase | Wavelength (Angstroms) |
|----------|-----------|------------|----------|------------|-------|------------------------|
| W Peg    | 23 19 51  | +26 16 44  | M7e      | 8.2        | 12.7  | 330.00                 |
| W Lyr    | 18 14 56  | 36 40 13   | M4.5e    | 7.9        | 12.2  | 195.50                 |
| RS Her   | 17 21 42  | +22 55 16  | M5e      | 7.9        | 12.5  | 217.68                 |
| W Her    | 16 35 12  | +37 20 42  | M3e      | 8.3        | 13.5  | 280.15                 |
| R Ser    | 15 50 42  | +15 08 03  | M7IIIe   | 5.7        | 14.4  | 357.74                 |
| S CrB    | 15 21 24  | +31 22 04  | M7e      | 6.5        | 14.1  | 367.23                 |
| V Boo    | 14 29 45  | +38 51 41  | M6e      | 7.0        | 11.3  | 258.95                 |
| R CVn    | 13 48 57  | +39 32 34  | M6IIIe   | 7.7        | 11.9  | 328.00                 |
| V CVn    | 13 19 27  | +45 31 38  | M6IIIa   | 6.8        | 8.8   | 188.74                 |
| R Tau    | 04 28 18  | +10 09 48  | M6e      | 8.6        | 14.2  | 313.34                 |
| R Tri    | 02 37 02  | +34 15 52  | M4IIIe   | 6.2        | 11.7  | 265.48                 |
| o W And  | 02 17 33  | +44 18 20  | M7:p     | 6.7        | 14.5  | 386.08                 |
| T Cas    | 00 23 14  | +55 47 34  | M7e      | 7.3        | 12.4  | 433.67                 |
| Cyg      | 19 50 33  | +32 54 53  | MS        | 5.2        | 13.4  | 419.97                 |

### Diagram

- Note: The diagram shows spectral lines and normalized flux measurements across different wavelengths.
- Specific periods and phases are marked with corresponding dates.

### Wavelength (Angstroms)

- The range varies from 5500 to 7500 Å.
- Important markers include Ca II and TiO absorption bands.
Table 2. Log of Observations

| Star  | Date (UT) | JD 2450000+ | Phase | V  | time (sec) | Telescope |
|-------|-----------|-------------|-------|----|------------|-----------|
| T Cas | 1 Nov 97  | 753.5       | 0.04  | 8.0| 120        | SARA      |
|       | 15 Nov 98 | 1132.5      | 0.92  | 10.2| 300        | SARA      |
| W And | 2 Nov 97  | 754.5       | 0.30  | 13.4| 900        | SARA      |
| o Cet | 1 Nov 97  | 753.5       | 0.80  | 7.1 | 60         | SARA      |
| R Tri | 15 Feb 97 | 494.5       | 0.73  | 10.0| 1800       | SARA      |
|       | 13 Mar 97 | 520.5       | 0.83  | 8.6 | 600        | DSO       |
|       | 1 Nov 97  | 753.5       | 0.70  | 10.3| 180        | SARA      |
| R Tau | 15 Feb 97 | 494.5       | 0.16  | 11.7| 360        | SARA      |
|       | 1 Nov 97  | 753.5       | 0.99  | 8.9 | 180        | SARA      |
|       | 14 Nov 98 | 1131.5      | 0.20  | 11.2| 300        | SARA      |
| U Ori | 14 Feb 97 | 493.5       | 0.30  | 10.7| 600        | SARA      |
|       | 1 Nov 97  | 753.5       | 0.01  | 6.6 | 60         | SARA      |
|       | 15 Feb 98 | 859.5       | 0.29  | 10.5| 600        | SARA      |
|       | 14 Nov 98 | 1131.5      | 0.03  | 6.4 | 300        | SARA      |
| R Leo | 14 Feb 97 | 493.5       | 0.16  | 8.5 | 600        | SARA      |
|       | 13 Mar 97 | 520.5       | 0.25  | 9.2 | 600        | DSO       |
|       | 28 May 97 | 596.5       | 0.49  | 10.0| 600        | SARA      |
|       | 15 Feb 98 | 859.5       | 0.34  | 9.8 | 300        | SARA      |
|       | 28 Mar 98 | 900.5       | 0.48  | 10.0| 300        | DSO       |
|       | 24 May 98 | 957.5       | 0.66  | 9.2 | 60         | SARA      |
|       | 14 Nov 98 | 1131.5      | 0.23  | 8.8 | 300        | SARA      |
| V CVn | 14 Feb 97 | 493.5       | 0.12  | 7.0 | 240        | SARA      |
|       | 13 Mar 97 | 520.5       | 0.27  | 7.9 | 600        | DSO       |
| Star   | Date (UT) | JD 2450000+ | Phase | V   | time (sec) | Telescope |
|--------|-----------|-------------|-------|-----|------------|-----------|
| 25 May 97 | 593.5     | 0.65        | 8.8   | 600 | SARA       |
| 15 Feb 98 | 859.5     | 0.06        | 7.1   | 600 | SARA       |
| 23 May 98 | 956.5     | 0.57        | 8.8   | 300 | SARA       |
| R CVn   | 15 Feb 97 | 494.5       | 0.09  | 8.5 | 600        | SARA      |
| 13 Mar 97 | 520.5     | 0.17        | 9.5   | 600 | DSO        |
| 26 May 97 | 583.5     | 0.39        | 11.3  | 1800| SARA       |
| 15 Feb 98 | 859.5     | 0.20        | 9.9   | 600 | SARA       |
| 23 May 98 | 956.5     | 0.50        | 11.9  | 300 | SARA       |
| V Boo   | 15 Feb 97 | 494.5       | 0.12  | 8.3 | 600        | SARA      |
| 26 May 97 | 595.5     | 0.51        | 11.3  | 600 | SARA       |
| 24 May 98 | 957.5     | 0.53        | 8.5   | 240 | SARA       |
| R Boo   | 25 May 97 | 583.5       | 0.11  | 8.2 | 600        | SARA      |
| S CrB   | 25 May 97 | 593.5       | 0.48  | 13.6| 1800       | SARA      |
| 24 May 98 | 957.5     | 0.47        | 13.5  | 600 | SARA       |
| R Ser   | 25 May 97 | 593.5       | 0.26  | 6.5 | 1200       | SARA      |
| 23 May 98 | 956.5     | 0.27        | 6.9   | 300 | SARA       |
| W Her   | 26 May 97 | 594.5       | 0.86  | 9.8 | 1200       | SARA      |
| 24 May 98 | 957.5     | 0.15        | 11.4  | 900 | SARA       |
| RS Her  | 28 May 97 | 596.5       | 0.75  | 11.0| 1800       | SARA      |
| 23 May 98 | 956.5     | 0.40        | 12.3  | 420 | SARA       |
| T Her   | 26 May 97 | 594.5       | 0.15  | 10.3| 1200       | SARA      |
| 24 May 98 | 957.5     | 0.37        | 12.4  | 1200| SARA       |
| W Lyr   | 26 May 97 | 594.5       | 0.94  | 8.9 | 900        | SARA      |
| Star   | Date (UT) | JD 2450000+ | Phase | V     | time (sec) | Telescope |
|--------|-----------|-------------|-------|-------|------------|-----------|
| χ Cyg | 26 May 97 | 594.5       | 0.66  | 11.8  | 300        | SARA      |
|        | 1 Nov 97  | 753.5       | 0.04  | 7.6   | 60         | SARA      |
|        | 19 Jul 98 | 1013.5      | 0.66  | 11.3  | 120        | SARA      |
| T Cep  | 2 Nov 97  | 754.5       | 0.52  | 10.2  | 300        | SARA      |
| W Peg  | 1 Nov 97  | 753.5       | 0.13  | 9.0   | 600        | SARA      |
Phase Dependent Spectroscopy of Mira Variable Stars

Michael W. Castelaz\textsuperscript{1,2}, Donald G. Luttermoser
Department of Physics, Box 70652, East Tennessee State University, Johnson City, TN 37614 and Southeastern Association for Research in Astronomy

Daniel B. Caton
Department of Physics and Astronomy, Dark Sky Observatory, Appalachian State University, Boone, NC 28608

Robert A. Piontek\textsuperscript{1,3}
University of Maryland, Department of Astronomy, College Park, MD 20742

ABSTRACT

Spectroscopic measurements of Mira variable stars, as a function of phase, probe the stellar atmospheres and underlying pulsation mechanisms. For example, measuring variations in TiO, VO, and ZrO with phase can be used to help determine whether these molecular species are produced in an extended region above the layers where Balmer line emission occurs or below this shocked region. Using the same methods, the Balmer-line increment, where the strongest Balmer line at phase zero is H\textgreek{d} and not H\textgreek{a} can be measured and explanations tested, along with another peculiarity, the absence of the H\textgreek{c} line in the spectra of Miras when the other Balmer lines are strong. We present new spectra covering the spectral range from 6200 Å to 9000 Å of 20 Mira variables. A relationship between variations in the Ca II IR triplet and H\textgreek{a} as a function of phase support the hypothesis that H\textgreek{c}’s observational characteristics result from an interaction of H\textgreek{c} photons with the Ca II H line. New periods and epochs of variability are also presented for each star.

Subject headings: variable stars: Miras, long period variables — low resolution spectroscopy: TiO, Balmer lines, Ca II IR triplet

\textsuperscript{1}Visiting Astronomer, Dark Sky Observatory, Appalachian State University

\textsuperscript{2}Also at The Pisgah Astronomical Research Institute

\textsuperscript{3}Participated in the Summer 1998 Southeastern Association for Research in Astronomy Research Experience for Undergraduates Program Sponsored by the National Science Foundation
1. Introduction

Mira-type variable stars are large, cool stars whose visual light variations exceed 2.5 magnitudes over periods from 150 days to \( \sim 500 \) days. The light curves of Mira variables depend on the surface temperature, radius, and opacity, all which vary as the star pulsates. These pulsations extend the atmosphere beyond that of the hydrostatic equilibrium configuration and enhance mass loss in these stars (Maciel 1977; Willson \& Hill 1979; Bertschinger \& Chevalier 1985; Bowen 1988; Fleischer et al. 1992, 1995). As a result, Mira variables are an important component in seeding the interstellar medium with C, N, and O.

These stars are located on the asymptotic giant branch, a transitional phase in stellar evolution. Photometric and spectroscopic measurements of their light curves provide a means to probe the stellar atmospheres and underlying pulsation mechanisms occurring during this stellar phase. In the near infrared, the spectra of Mira stars are dominated by the TiO \( \gamma \) system, the VO \( \gamma \) system and ZrO (Wing 1967). The TiO features are thought to be produced in a layer somewhat far from the photosphere (Gillet 1988). Haniff et al. (1992) present optical aperture synthetic images of the photosphere of \( o \) Ceti at 6500 Å, 7007 Å, and within a TiO bandhead at 7099 Å, with the star phase \( \sim 0.94 \). They find asymmetry in the images, with the TiO image one and a half times larger than the photospheric images. Also, narrowband speckle interferometric measurements taken in the TiO 7120 Å bandhead and outside at 7400 Å by Labeyrie et al. (1977) shows that the diameters of R Leo and \( o \) Ceti are twice as large in the TiO feature than outside of it. This demonstrates that a model atmosphere, based on the spectra observed over a TiO bandpass, provides parameters such as \( T_{\text{eff}} \) and \( \log(g) \) in an atmospheric layer far from the photosphere.

Joy (1926) presents a comprehensive phase dependent spectroscopic (35 Å/mm) study of a prototype Mira variable, \( o \) Ceti. Analyzing 131 spectra taken over a ten year period, Joy describes several important characteristics. Briefly, the spectra of \( o \) Ceti show that TiO bands vary with magnitude, hydrogen emission lines appear with greatest intensity at or shortly after maximum visual brightness (phase zero), ionized iron emission lines are observed at maximum, and the temperature is estimated to vary from 1800 K to 2300 K. Absorption lines (including iron, vanadium, chromium, manganese, calcium, and magnesium) were used to measure a variation in radial velocity. The maximum positive velocity occurs at phase zero, and greatest blueshift at minimum light. Later, Joy (1954) took 88 spectrograms of \( o \) Ceti at a higher spectral resolution (typically 10.3 Å/mm) and confirmed that maximum velocity of recession occurs soon after visual maximum. He attributed these results to a pulsational mechanism, and suggested the possibility of shocks.

Perhaps one of the most interesting characteristics of Mira spectra is the strong hydrogen Balmer line emission that is seen throughout much of the pulsation cycle. As
Pickering (1887) first noticed in spectra of \(\omega\) Ceti, and later described in detail by Joy (1926, 1947, 1954), the hydrogen Balmer emission line flux in relation to the nearby photospheric (i.e., pseudocontinuum) flux is unique in the oxygen-rich (M-type) Mira spectra: Balmer \(\text{H}_\alpha\) emission is typically weaker than \(\text{H}_\beta\) which in turn is weaker than \(\text{H}_\gamma\) near peak visual brightness. \(\text{H}_\delta\) is seen as the strongest Balmer emission line at phase zero. Lines higher in the series (i.e., towards shorter wavelengths), are weaker. This Balmer increment (i.e., \(F_{\text{H}_\alpha} < F_{\text{H}_\beta} < F_{\text{H}_\gamma} < F_{\text{H}_\delta}\)) is just opposite of what would be expected, \(\text{H}_\alpha\) having the largest oscillator strength, should be stronger than \(\text{H}_\beta\) and the higher order Balmer lines should be weaker down the line (i.e., one should see a Balmer line decrement), assuming these lines all form in the same region of the atmosphere (i.e., similar \(T\) and \(P\)). Meanwhile, in S-type and carbon-star (N-type) Miras, the strength of the Balmer lines approximately follow their expected respective oscillator strengths (Merrill 1940).

For years, this Balmer-line increment in the M-type Miras has been attributed to TiO absorption which hide \(\text{H}_\alpha\), \(\text{H}_\beta\), and \(\text{H}_\gamma\) fluxes (Merrill 1940; Joy 1947; Gillet 1988), although there is some debate of the extent that this or other molecular absorption has on the \(\text{H}_\alpha\) line (Gillet, Maurice, & Baade 1983; Gillet et al. 1985). Recently another explanation has been given for this Balmer increment: NLTE radiative transfer calculations of hydrodynamic models representative of Mira variables (Bowen 1988) suggest that the Balmer increment results from radiative transfer effects in the hydrogen lines themselves when formed in a shocked atmosphere (Luttermoser & Bowen 1992; Luttermoser, Bowen, & Willson 2000, in preparation). In these calculations, \(\text{H}_\alpha\), having the highest optical depth, forms just in front of the innermost shock. \(\text{H}_\beta\) then forms a little deeper, due to its lower optical depth, and \(\text{H}_\gamma\) deeper still. The optical depth of \(\text{H}_\delta\) causes it to arise from the hottest part of the shock, hence extends higher above the continuum then the longward Balmer lines. Then as one goes to higher-order lines in the series, the opacity in these lines is not high enough in the shock for these lines to form there — we see through the shock at these transitions. As such, this increment may be giving us information of the shock thickness for Miras and may indicate that the shock structure of the S-type and N-type Miras are fundamentally different than the M-type Miras. Future NLTE radiative hydrodynamic models of Miras with different C/O ratios are needed to see if this is the case. It is likely that a combination of both processes (i.e., TiO absorption and NLTE radiative transfer effects) are responsible for this Balmer line increment.

Another striking feature is the weakness (and often absence) of the \(\text{He}\) line (3970.074 Å) in the Balmer series near maximum visual brightness (see Gillet 1988). Merrill (1940) noted this and suggested this weakness a result of the interaction between the \(\text{He}\) transition and the \(\text{Ca II H}\) line (3968.470 Å) wing. Castelaz & Luttermoser (1997, hereafter CL97) concur with this suggestion. Briefly, the \(\text{He}\) photons may be scattered by the \(\text{Ca II H}\) line out to
IR wavelengths via the 8662 Å line. Of the three lines in the Ca II IR triplet, the 8498 Å and 8542 Å lines share the same upper level of the Ca II K line, whereas the 8662 Å line shares the same upper level as the Ca II H line. As such, if He photons are being scattered by Ca II H, the 8662 Å will show variations independent of the other two Ca II IR triplet lines and can be tested by monitoring the strength of the absorption of this 8662 Å line as compared to the 8498 Å and 8542 Å lines as a function of phase in the Mira stars.

Ca II IR triplet observations of Miras have been carried out by Contadakis & Solf (1981) and Gillet et al. (1985). Unfortunately, Contadakis & Solf (1981) only made one observation of the $\lambda$8662 line in their monitoring program of S-type Miras. They observed the $\lambda$8498 and $\lambda$8542 in emission in many of the stars in their sample near phase 0 and as absorption lines at other phases. So, we cannot use their data to test our proposed hypothesis. Gillet et al. (1985) observed P Cygni profiles in the Ca II IR triplet lines near phase 0 in the hot Mira variable, S Car. Once again, only one spectrum of the 8662 Å line was obtained in their sample of Ca II spectra. CL97 present a set of phase dependent spectra taken especially to address this problem of He photon scattering by Ca II. Spectra of seven Mira variables were taken at different phases and suggest a possible anticorrelation between H$\alpha$ emission and Ca II $\lambda$8662 absorption. Assuming that the He line strength variations are in phase with H$\alpha$, then the apparent anticorrelation between the strength of the H$\alpha$ emission line and the strength of the Ca II $\lambda$8662 line suggests that a fluorescence is taking place in the Ca II $\lambda$8662 line with He$\epsilon$ serving as the pump through the Ca II H line. In this paper we present phase dependent spectra of twenty Mira variables to further explore the anticorrelation of Ca II $\lambda$8662 absorption with H$\alpha$ emission.

2. Observations

The twenty stars for which spectra are presented, their equatorial coordinates, mean spectral types, and visual maxima and minima, taken from the SIMBAD database are presented in Table 1. Also in Table 1 are new ephemerides of these Mira stars. We calculated the ephemerides from AAVSO light curves measured within the previous eight years. The curves were fit with a linear combination of sine and cosine functions from which a new period and JDO were determined for each star.

Spectra of Mira variable stars taken between February 1997 and November 1998 using a low resolution spectrograph. Since we are interested in only monitoring absorption and emission line strengths for this program, low-resolution spectra are all that is required. The spectrograph was used at both the Southeastern Association for Research in Astronomy (SARA) 0.9-m telescope at Kitt Peak, and Appalachian State University’s Dark Sky
Observatory (DSO) 0.45-m telescope located near Boone, NC. A converter lens was used at both sites to convert the respective telescope f-ratio to about f/11 for the spectrograph.

The spectrograph was configured with a 600 g/mm grating. The slit width was 100 µm. At the SARA 0.9-m telescope, the slit width is 3 arcseconds, and at the DSO 0.45-m telescope, the slit width is 4 arcseconds. At both sites, the slit was parallel to the hour angle. A cooled 768 x 512 CCD camera with 9 µm square pixels was used to record the spectra. The linear dispersion is 1.08 Å per pixel, covering 768 pixels, or 829 Å on the CCD. The spectral resolution was measured to be 2.4 Å. By rotating the grating up to four times per star, spectra were taken from about 6200 Å to 9000 Å and include Hα, TiO, VO, and the Ca II infrared triplet lines. Integration times were adjusted to achieve a signal-to-noise greater than 100 for most spectra, except for the May 1997 spectra of W Her, T Her, and W Lyr, and the May 1998 spectrum of T Her.

Table 2 gives the log of observations for each star, which includes dates of observation, phase of the variable, approximate visual magnitude, integration time, and observing site. The phases listed in Table 2 were determined from the ephemerides given in Table 1, and refer to the visual phases, with phase zero corresponding to maximum visual brightness. The visual magnitudes have been obtained from curve fits to the light curves from the AAVSO database.

Dark frames and flat frames were taken for flat fielding purposes. Spectra of neon lamp emission were taken simultaneously with the stellar spectra and used for wavelength calibration. We flat-fielded the images and extracted the spectra using MIRA software. The extracted spectra were wavelength calibrated using the spectrum of neon superimposed on the CCD frame with the stellar spectrum (Crowe, Heaton, & Castelaz 1996).

3. Results

The spectra for the Mira variable stars are shown in Figure 1. The wavelengths of TiO, VO, ZrO, the Ca II IR triplet, Hα, and terrestrial oxygen are marked above each set of spectra. Their wavelengths were identified in CL97. The ZrO and VO absorption overlap at 6574 Å and 6578 Å, and VO and TiO absorb at 7865 Å and 7861 Å, respectively. Due to the low dispersion of our spectra, these features are blended. The appearance of a relatively narrow feature near 8230 Å is seen in some of the spectra - an o Cet spectrum taken 1 November 1997 (phase 0.80), and spectra taken 26-28 May 1997 of R Leo (0.49), V CVn (0.65), R Boo (0.11), R Ser (0.26), RS Her (0.75), and W Lyr (0.94). This feature is due to terrestrial H$_2$O at 8227 Å (Turnshek et al. 1985) and is an effective measure of the
relative humidity in the air.

### 3.1. Radiative Transfer in the Ca II Ion

The transfer of radiation in Ca II ion is very complicated. Beside He photons affecting the level populations of the 3d $^{2}P_{1/2}$ state through the Ca II H line, the hydrogen Lyman-α line lies just shortward of the Ca II ionization edge of the 3d $^{2}D_{3/2}$ (at 1218.1 Å) and the 3d $^{2}D_{5/2}$ (at 1219.0 Å) states. These two states are the lower levels of the Ca II IR triplet lines and are metastable. If Lyman-α is a strong emission feature, and the Ca II $^{2}D$ continuous opacity forms in a region of the atmosphere where the Lyman-α line is not yet in detailed balance, then Lyman-α photons may influence the Ca II IR triplet lines as well. Note that no observations have yet been made in the far-UV for Miras.

Line center of Lyman-α lies 2.4 Å shortward of the $^{2}D_{3/2}$ ionization edge and 3.3 Å shortward of the $^{2}D_{5/2}$ edge. So, much of the Lyman-α emission profile can ionize Ca II out of the metastable state. The lower level of the Ca II 8662 Å line transition is the $^{2}D_{3/2}$ state whose ionization edge lies slightly closer to Lyman-α than the $^{2}D_{5/2}$ edge. The following question arises: Will Lyman-α photons affect the level densities of the two 3d $^{2}D$ states differently? To answer this question, we ran a few atmospheric models with an arbitrarily located 10,000 K shock through the LTE stellar atmosphere code ATLAS (Kurucz 1970; Brown et al. 1989). Although it is likely that NLTE effects will dominate the level and ion densities in the atomic and molecular species (Luttermoser & Bowen 1992; Luttermoser, Bowen, & Willson 2000, in preparation), these LTE runs are performed to merely determine the variation of the Ca II $^{2}D$ continuous opacity across the Lyman-α profile. It also should be pointed out that in regions of the atmosphere where LTE no longer applies, the assumption of radiative and hydrostatic equilibria are no longer valid either in these pulsating giant stars. In fact, a Mira star has numerous shocks propagating through its atmosphere at any given time as has been shown by Bowen (1988) and more recently by Höfner, et al. (1998) and Loidl, et al. (1999). Willson (2000) gives a very detailed review of all the dynamic modeling that has been performed on these pulsating stars and discusses the problems of carrying out NLTE radiative transfer in such a dynamic atmosphere.

We sampled atmospheric depths in front of the shock, in the shock, behind the shock, and deep in the photosphere where the continuum reaches optical depth unity in this region of the spectrum. We found that the continuous opacity from the Ca II 3d $^{2}D$ ionization remained constant (from both $J$ sub-levels) to within 0.15% from the location of the edges through 1210 Å, which should include most of the Lyman-α emission profile. The fact that the continuous opacity from Ca II remains constant across the Lyman-α profile
indicates that Lyman-\(\alpha\) will not preferentially affect the number density in the \(^2\!D_{3/2}\) level as compared to the \(^2\!D_{5/2}\) level — photoionizations of Ca II due to Lyman-\(\alpha\) photons will affect the strengths of three Ca II IR lines in a similar fashion. Therefore, any variation in the 8662 Å line that is not seen in the other lines must result from some process other than Lyman-\(\alpha\) photoionizations.

3.2. The H\(\alpha\) Emission Line and Ca II IR Triplet

We are interested in the strength of the Ca II IR triplet compared to the strength of H\(\alpha\) as a function of phase, since we are assuming that variations in H\(\alpha\) will mimic variations in H\(\alpha\). The Ca II IR triplet lines are not strong, as expected for stars later than M0 (Zhou 1991). A total of twenty-seven spectra of fifteen stars in our sample span the wavelength range from H\(\alpha\) through the Ca II IR triplet. The remaining spectra are missing either the H\(\alpha\) or the Ca II IR triplet regions of the spectrum because the spectrograph grating was not rotated sufficiently during observation to cover those parts of the spectrum.

The observations which show obvious H\(\alpha\) emission features include U Ori (phase 0.29), R Leo (0.34), V CVn (0.12), R CVn (0.09 and 0.17), R Ser (0.26), W Lyr (0.94), and \(\chi\) Cyg (0.04). At the same phases, the Ca II \(\lambda8662\) is seen in emission in R Leo, R Ser, W Lyr, and \(\chi\) Cyg, whereas the \(\lambda8498\) and \(\lambda8542\) lines stay in absorption or are not apparent. Ca II \(\lambda8662\) is not seen in emission in any other spectra that we took, only in those that show H\(\alpha\) emission. The Ca II IR triplet can be seen in absorption in the remaining twenty-three spectra (although in some cases weakly), except for U Ori (phase 0.01), S CrB (0.48), W Her (0.86), and T Her (0.15) which do not appear to have any type of Ca II IR triplet features.

At this point, we call attention to three of the hydrogen Paschen lines which lie close to each Ca II IR triplet line: 8502.4 Å (4.4 Å redward from the Ca II \(\lambda8498.0\) line), 8545.3 Å (2.8 Å redward of the \(\lambda8542.1\) line), and 8665.0 Å (2.8 Å from the \(\lambda8662.1\) line), the Pa13, Pa12, and Pa10 lines respectively. It has been shown by Gillet et al. (1985) that even though the Pa-\(\delta\) line is in emission in the spectrum of the Mira variable S Car, the higher order Paschen lines near the Ca II IR triplet are neither seen in absorption nor emission, analogous to the Balmer lines. As such, it is unlikely that these higher order Paschen lines are affecting the Ca II lines in our spectra.
3.3. The TiO Bands and H\(\alpha\) Emission

In addition to the Ca II IR triplet, we are interested in comparing the TiO \(\gamma\) system molecular features to H\(\alpha\) emission as a function of phase. A qualitative comparison of H\(\alpha\) emission with the molecular features can be made for U Ori, R Leo, R CVn, R Ser, W Lyr, and \(\chi\) Cyg, stars with phase dependent spectra that also show the H\(\alpha\) emission line in at least one spectrum. Weak H\(\alpha\) emission is seen in U Ori at phase 0.29 on 15 Feb 1998. However on the previous pulsation cycle H\(\alpha\) emission is not obvious at either phases 0.30 or 0.01, although there may some very weak emission. Stronger H\(\alpha\) emission is seen in R Leo (phase 0.34), R CVn (0.09, 0.17, and 0.39), R Ser (0.26), W Lyr (0.94), and \(\chi\) Cyg (0.04).

As mentioned earlier in the paper, H\(\alpha\) is notorious for being observed as a weak emission feature when H\(\beta\), H\(\gamma\), and H\(\delta\) are strong. Surprisingly though, H\(\alpha\) emission was not seen in various spectra where we would expect to find it: R Tau (phase 0.16 and 0.99), R Leo (0.16 and 0.25), V CVn (0.12, 0.27, and 0.06), R CVn (0.20), V Boo (0.12), and T Her (0.15).

Merrill (1940) and Joy (1926) report that the TiO bands are regularly stronger at minimum than at maximum light in Mira variables. To measure this trend, we checked the variability of the TiO feature at 7054 Å (bandhead) with respect to a portion of the flux uncompromised by TiO, VO, and ZrO. We integrated the flux of each spectrum in the 6995-7045 Å wavelength band (non-TiO) and the 7060-7110 Å band (TiO), each 50 Å wide. We then divided the integrated flux of the TiO band by the non-TiO band. By doing this, any scattered light that may be in the spectra are effectively canceled out. Since our data set contains warm oxygen-rich (earlier than M6), cool oxygen-rich (M6-M8), and one MS Mira sampled sporadically over various phases, we only include the cool oxygen-rich Miras in Figure 2, which graphs the above mentioned flux ratio as a function of phase. The M6-M8 stars are selected here in order to minimize the spread of effective temperatures at maximum light which will influence the strength of the TiO bands. Ideally, one would want many observations of each star over a single pulsation cycle. However, with the limited sample we have, we feel that we can get an approximate test of TiO variation with respect to phase. There appears to be no obvious trend in variations in the TiO 7054 Å band flux with respect to the non-TiO band flux. Due to this observation, variations seen in H\(\alpha\) as a function of phase must primarily result from variations in the H\(\alpha\) emission itself and not to variations in overlying TiO absorption. As a result, using H\(\alpha\) flux variations as a proxy to variations in the intrinsic He flux is valid from this analysis.
4. Discussion

We wish to test the idea that the apparent lack of Hǫ emission at 3970 Å when the other Balmer lines are strong emission features is anticorrelated with the strength of the Ca II absorption line at 8662 Å. As reported in the Introduction, this anticorrelation results from Hǫ photons being scattered by the Ca II H line out to the Ca II line at 8662 Å, causing this Ca II absorption line to be filled in with respect to the other two Ca II IR triplet lines.

We use Hα as a proxy for the Hǫ line. Following the same analysis presented by CL97, we determined a relative line strength, $F$, for Hα emission and the Ca II IR triplet absorption lines. Two points were selected on either side of the emission or absorption feature. The wavelengths of these points were kept constant for all measurements. The observed profile was integrated across the wavelength window defined by these two points resulting in an integrated flux $f_e$. A straight line connected between these two points represent a pseudocontinuum and the integrated flux, $f_c$, calculated for it. Then

$$ F = \frac{f_e - f_c}{f_c}, $$

where $F$ will be negative for absorption lines and positive for emission lines. The measurements were done for the stars in our sample where their spectra included wavelengths below 6563 Å and above 8662 Å; a total of 27 spectra. Figure 3 shows the relative line strengths of Hα, and the Ca II λ8498, Ca II λ8542, and Ca II λ8662 absorption lines as a function of phase. The uncertainty in the measurements is about ±0.007. The relative line strength of Hα clearly shows large scatter near visible maximum, and is zero within the uncertainty from phase 0.5 to 0.8. This is consistent with Balmer emission lines becoming prominent near maximum visible light. The Ca II λ8498 line strengths are zero, within the uncertainty of the measurements; variation is not observed in this line. The Ca II λ8542 line does show some scatter near phase 0.0, and is zero after phase 0.2. The Ca II λ8662 line shows more scatter than either of the other two Ca II IR triplet lines, particularly near phase 0.0.

From Figure 3, it is difficult to see any correlation between the relative line strength of the Hα and Ca II IR triplet lines. However, we can plot the relative line strengths of Ca II IR triplet lines versus the strength of the Hα line to look for correlations. Figure 4 shows the relative line strengths of Ca II λ8498, Ca II λ8542, and Ca II λ8662 versus the relative line strength of Hα. The data of each plot is linearly fit and the results of the fits are drawn in the plots. The slopes of the Ca II λ8498 and λ8542 relative line strengths versus Hα are 0.08 and 0.05, respectively. The slope of the linear fit of the Ca II λ8662 relative line strength versus Hα relative line strength, on the other hand, is 0.32, which is
significantly different than the other two Ca II IR triplet linear fits. Furthermore, the Ca II λ8662 versus Hα relative line strength slope is positive, so that as the strength of the Hα line increases, the strength of the Ca II λ8662 line decreases (i.e. becomes more positive) and even goes into emission. This is the effect we expect if He photons are being scattered by the Ca II H line out to the Ca II line at 8662 Å, causing the Ca II λ8662 absorption line to be filled in with respect to the other two Ca II IR triplet lines.

5. Conclusion

The 6200 Å to 9000 Å spectra of Mira variables taken at different phases support a possible anticorrelation between Hα emission and Ca II λ8662 absorption as first suggested by CL97. Assuming that the He line strength variations are in phase with Hα, then the apparent anticorrelation between the strength of the Hα emission line and the strength of the Ca II λ8662 line suggests that a fluorescence is taking place in the Ca II λ8662 line with He serving as the pump through the Ca II H line. This type of fluorescence is common in Mira type variables. The strong Fe I (42) lines at 4202 Å and 4308 Å seen in Miras are well known fluoresced features; in this case, the ultraviolet Mg II h & k lines serve as the pump via an Fe I (UV3) transition (e.g., Bidelman & Herbig 1958; Willson 1976; Luttermoser 1996).

The next phase of this research program is to systematically determine $T_{\text{eff}}$ and $\log g$ as a function of phase for the Miras in our sample. A detailed description of the LTE model synthetic spectra is given by Piontek & Luttermoser (1999).

ACKNOWLEDGEMENTS. MWC greatly appreciates support from NSF Grant AST-9500756 which was the primary source of support this research. The long-term observations required for this research project would not be possible without the continued commitment of the Southeastern Association for Research in Astronomy to provide the periodic observing times on a meter-class telescope at Kitt Peak. We thank Dr. Peter Mack of Astronomical Consultants and Equipment for his excellent technical expertise at the SARA Observatory. We also thank Robert Miller at Appalachian State University who machined the spectrograph adapter so we could use the spectrograph on the DSO 0.45-m telescope. We are indebted to Marie Rinkoski, NSF sponsored SARA REU student during the Summer 2000, who kindly calculated the phases while reducing a new set of spectra as part of her research. We also greatly appreciate the comments of the referee of this paper, resulting in several significant improvements. In this research, we have used, and acknowledge with thanks, data from the AAVSO International Database, based on observations submitted to the AAVSO by variable star observers worldwide. This research
has made use of the SIMBAD databases, operated at CDS, Strasbourg, France.

6. References

Bessell, M.S., Scholz, M. & Wood, P.R. 1996, A& A, 307, 481

Bidelman, W.P. & Herbig, G.H. 1958, PASP, 70, 451

Bowen, G.H. 1988, ApJ, 329, 299

Brown, J.A., Johnson, H.R., Alexander, D.R., Cutright, L., Sharp, C.M. 1989, ApJS, 71, 623

Castelaz, M. W. & Luttermoser, D. G. 1997, AJ, 114, 1584

Celis S., L. 1984, AJ, 89, 527

Crowe, K., Heaton, B., Castelaz, M. W. 1996, IAPPP Communications, 68, 30

Contadakis, M. E., & Solf, J. 1981, A& A, 101, 241

Gillet, D. 1988, A&A, 192, 206

Gillet, D., Maurice, E., & Baade, D. 1983, A& A, 128, 384

Gillet, D., Maurice, E., Bouchet, P., & Ferlet, R. 1985, A& A, 148, 148

Haniff, C. A., Ghez, A. M., Gorham, P. W., Kulkarni, S. R., Matthews, K., & Neugebauer, G. 1992, AJ, 103, 1662

Höfner, S., Jørgenson, U.G., Loidl, R., & Aringer, B. 1998, A&A, 340, 497

Joy, A. H. 1926, ApJ, 63, 281

Joy, A. H. 1947, ApJ, 106, 288

Joy, A. H. 1954, ApJS, 1, 39

Kurucz, R. L. 1970, Smithsonian Ap. Obs. Special Rept., No. 309

Labeyrie, A., Koechlin, L., Bonneau, D., Blazit, A., & Foy, R. 1977, ApJ Letters, 218, L75

Lockwood, G. W. 1973, ApJ, 180, 845

Lockwood, G. W. & Wing, R. F. 1971, ApJ, 169, 63
Loidl, R., Höfner, S., Jørgenson, U.G., & Aringer, B. 1999, A&A, 342, 531

Luttermoser, D.G. 1996, in Ninth Cambridge Workshop on Cool Stars, Stellar Systems, and the Sun, R. Pallavicini & A.K. Dupree, ASP Conference Series, 109, 535

Luttermoser, D. G. & Bowen G. H. 1992, in Seventh Cambridge Workshop on Cool Stars, Stellar Systems, and the Sun, M.S. Giampapa & J.A. Bookbinder, ASP Conference Series, 26, 558

Mattei, J. A., 1996, AAVSO Bulletin 59, Predicted Dates of Maxima and Minima of Long Period Variables for 1996

Merrill, P. W. 1940, in Spectra of Long-Period Variable Stars, (Univ. of Chicago Press; Chicago), p. 44

Pickering, E. C. 1887, Nature, 36, 32

Piontek, R. & Luttermoser, D. L. 2000, in Bull. of the AAS, 31, 1238

Turnshek, D. E., Turnshek, D. A., Graine, E. R. & Boeshaar, P.C. 1985, An Atlas of Digital Spectra of Cool Stars, (Western Research Corp.; Tucson, AZ USA)

Willson, L. A. 1976, ApJ, 205, 172

Willson, L.A. 2000, Annu. Rev. Astron. Astrophys., 38, 573

Wing, R. F. 1967, Ph.D. Thesis, University of California, Berkley

Zhou, X. 1991, A&A, 248, 367
Fig. 1.— Spectra of the 20 Mira variables. The name of the star, date, and phase, are given on each spectrum. The flux is normalized to one. Above each set of spectra are markers for the major spectral features, and terrestrial oxygen. H\(\alpha\) is weak, or not seen in most of the spectra. Markers enclose regions of individual spectra where no data was taken. Spikes due to cosmic ray hits have not been removed.

Fig. 2.— The integrated flux ratio of the 7060-7110 Å wavelength band (A) to that of the 6995-7045 Å band (B) plotted as a function of phase. Band A contains TiO opacity whereas band B is free of TiO, VO, and ZrO bandheads. This plot contains only data from the M6, M7, and M8 spectral-type Miras. No apparent trends are seen in the data, which indicates that variations in the H\(\alpha\) flux result primarily from intrinsic flux variations in H\(\alpha\) and not from varying overlying TiO absorption.

Fig. 3.— Relative line strengths of H\(\alpha\) (filled square), Ca II lines at 8498 Å (circle), 8542 Å (triangle), and 8662 Å (diamond), as a function of light-variation phase. Variation in the Ca II lines at 8498 Å and 8542 Å mimic each other, whereas the Ca II line at 8662 Å does not follow the same trend. The uncertainty of the measurements is ±0.007.

Fig. 4.— The relative line strengths of Ca II λ8498, Ca II λ8542, and Ca II λ8662 versus the relative line strength of H\(\alpha\). The data of each plot is linearly fit and the results of the fits are drawn in the plots. Only the Ca II λ8662 versus H\(\alpha\) plot shows a significant slope which implies a correlation between the occurrence of the two features.
CaII 8498
CaII 8542
CaII 8662

Hα