VLT/GIRAFFE spectroscopic observations of the metal-poor blue compact dwarf galaxy SBS 0335–052E

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ABSTRACT

Aims. We present two-dimensional spectroscopy of the extremely metal-deficient blue compact dwarf (BCD) galaxy SBS 0335–052E aiming to study physical conditions, element abundances and kinematical properties of the ionised gas in this galaxy.

Methods. Observations were obtained in the spectral range \( \lambda 3620 – 9400 \) Å with the imaging spectrograph GIRAFFE installed on the UT2 of the Very Large Telescope (VLT). These observations are the first ones carried out so far with GIRAFFE in the ARGUS mode which allows to obtain simultaneously 308 spectra covering a 11’’4 × 7’’3 region.

Results. We produced images of SBS 0335–052E in the continuum and in emission lines of different stages of excitation. We find that while the maximum of emission in the majority of lines, including the strong lines H\( \beta \) 4861 Å, H\( \alpha \) 6563 Å, [O iii] 4363,5007 Å, [O ii] 3726,3729 Å, coincides with the youngest south-eastern star clusters 1 and 2, the emission of He II \( \lambda 4686 \) Å line is offset to the more evolved north-west clusters 4, 5. This suggests that hard ionising radiation responsible for the He II \( \lambda 4686 \) Å emission is not related to the most massive youngest stars, but rather is connected with fast radiative shocks. This conclusion is supported by the kinematical properties of the ionised gas from the different emission lines as the velocity dispersion in the He II \( \lambda 4686 \) Å line is systematically higher, by \(~ 50\% – 100\%\), than that in other lines. The variations of the emission line profiles suggest the presence of an ionised gas outflow in the direction perpendicular to the galaxy disk. We find a relatively high electron number density \( N_e \) of several hundred \( \text{cm}^{-3} \) in the brightest part of SBS 0335–052E. There is a small gradient of the electron temperature \( T_e \) and oxygen abundance from the East to the West with systematically higher \( T_e \) and lower \( 12 + \log O/H \) in the western part of the galaxy. The oxygen abundances for the whole H II region and its brightest part are \( 12 + \log O/H = 7.29 \pm 0.02 \) and \( 7.31 \pm 0.01 \), respectively. We derive the He mass fraction taking into account all systematic effects. The He mass fraction \( Y = 0.2463 \pm 0.0030 \), derived from the emission of the whole H II region, is consistent with the primordial value predicted by the standard big bang nucleosynthesis model. We confirm the presence of Wolf-Rayet stars in the cluster 3.

Key words. galaxies: fundamental parameters – galaxies: starburst – galaxies: ISM – galaxies: abundances – galaxies: individual (SBS 0335–052E)

1. Introduction

The blue compact dwarf (BCD) galaxy SBS 0335–052E is an excellent nearby laboratory for studying star formation in low-metallicity environments. Since its discovery as one of the most metal-deficient star-forming galaxies known [Izotov et al. 1990], with oxygen abundance \( 12 + \log O/H \sim 7.30 \) [Melnick et al. 1992, Izotov et al. 1997b, 1999, Thuan & Izotov 2003], SBS 0335–052E has often been proposed as a nearby young dwarf galaxy [Izotov et al. 1990, 1997b, Thuan et al. 1997, Papaderos et al. 1998, Pustilnik et al. 2004]. Thuan et al. (1997) and Papaderos et al. (1998), using the same Hubble Space Telescope (HST) images, have found several luminous clusters. The brightest clusters are labelled in Fig. 1 which represents the highest spatial resolution archival UV HST/Advanced Camera for Surveys (ACS) image of SBS 0335–052E.
SBS 0335–052E

Fig. 1. Archival HST/ACS UV image of SBS 0335–052E with the labelled compact clusters. North is up and East is to the left.

obtained by Kunth et al. (2003). Some of the clusters are very young and produce extended regions of ionised gas (Melnick et al. 1992; Izotov et al. 1997b; Papaderos et al. 1998; Pustilnik et al. 2004). In particular, Izotov et al. (2001b), using deep long-slit spectra of SBS 0335–052E, have shown that extended Hα emission is detected over ∼ 6–8 kpc, suggesting that hot ionised gas is spread out far away from the central part of the galaxy.

Thuan & Izotov (1997) using HST/GHRS UV spectrum of SBS 0335–052E have discovered a very broad Lyα line in absorption suggesting that this galaxy is embedded in a large envelope of neutral gas. The column density of N(H i) = 7 × 10^22 cm^−2 in SBS 0335–052E derived by Thuan & Izotov (1997) is the largest one known for the BCDs. Later, Pustilnik et al. (2001) using Very Large Array (VLA) observations in the line H i λ21 cm have detected a large neutral gas cloud around SBS 0335–052E with a size 66 by 22 kpc elongated in the east-west direction and with two maxima separated by 22 kpc. The first maximum in H i distribution is connected to SBS 0335–052E, and the second one to the companion dwarf galaxy SBS 0335–052W discovered by Pustilnik et al. (1997). The latter galaxy is shown by Izotov et al. (2005) to be the lowest metallicity emission-line galaxy known with 12 + log O/H = 7.12 ± 0.03 (Thuan et al. 2003) using Far Ultraviolet Spectroscopic Explorer (FUSE) observations. Izotov et al. (1997b) and Thuan et al. (1997) have found variations of extinction in this galaxy from the optical spectroscopic and photometric observations. Later, Thuan et al. (1999) and Houck et al. (2004) using Infrared Space Observatory (ISO) and Spitzer mid-infrared observations have found an emission of an appreciable amount of warm dust with a characteristic temperature of ∼ 100K. Even hotter dust with a temperature of several hundred degrees is expected to be present in SBS 0335–052E which is indicated in the infrared spectra at shorter wavelengths, namely 2–4 μm (Vanzi et al. 2000; Hunt et al. 2004).

Since the discovery paper by Izotov et al. (1990), it is known that the ionising radiation in SBS 0335–052E is hard. The presence of the high-ionisation He ii λ4686Å emission line (e.g. Izotov et al. 1990; Melnick et al. 1992; Izotov et al. 1997b) and of the [Fe v] λ4227Å emission line (Izotov et al. 2001b; Frick et al. 2001) suggests that radiation with photon energies greater than 4 Rydberg is intense and could not be explained by stellar emission. Furthermore, Izotov et al. (2001b) and Thuan & Izotov (2005) have found [Fe v] and [Ne v] emission lines which require intense radiation with photon energies above 7 Rydberg. SBS 0335–052E has also been detected in the X-ray range (Thuan et al. 2004). The origin of the hard ionising radiation remains unclear. Several mechanisms such as radiation from most massive main-sequence stars, Wolf-Rayet stars, high-mass X-ray binaries and radiative shocks have been discussed by e.g. Garnett et al. (1991), Schaerer (1996), Izotov et al. (2001b), Izotov et al. (2004) and Thuan & Izotov (2005). The most recent investigations have shown that although the stellar origin of hard radiation is not completely excluded, the most likely source of hard radiation is fast radiative shocks.

Despite the efforts of different groups in studying the properties of SBS 0335–052E and its evolutionary status many problems remain unsolved. Since this galaxy is possibly a young galaxy it could be considered as a local counterpart of the high-redshift young dwarf galaxies. Therefore the continuation of its studies is important for cosmological applications. In this paper we present a two-dimensional spectroscopic study of SBS 0335–052E with the VLT/GIRAFFE. These are the first observations carried out so far in the ARGUS mode. Two new features are characteristic for these new observations which were not present in all previous spectroscopic studies of this galaxy. First, new observations allow to map the whole galaxy in different emission lines and in the continuum. This gives integrated characteristics of different emission lines for the whole H ii region, such as the line luminosities, which are necessary input parameters for building up the model of the H ii region. Second, the spectral resolution of new observations is by one order of magnitude better than in all previous spectroscopic observations of SBS 0335–052E and it is enough to obtain the intrinsic profiles of emission lines. This allows to study the kinematics of the H ii region, and to make a comparison of kinematic characteristics in regions of different ionisation stages. In particular, it is important to compare the He ii λ4686Å line profiles with the profiles of other emission lines and, hence, to make conclusions concerning the origin of hard radiation in SBS 0335–052E.
In this work, we describe the observations and data reduction. Morphology of SBS 0335–052E in different emission lines and continuum is considered in §3. Kinematic properties are discussed in §4. Heavy element abundances and helium abundance are derived in §5. The Wolf-Rayet stellar population in SBS 0335–052E is discussed in §6. Our conclusions are summarized in §7.

2. Observations and Data Reduction

Observations of SBS 0335–052E with the VLT/GIRAFFE spectrograph have been done during the nights 19–22 November, 2003 in the entire visible range. GIRAFFE is equipped with a 2K×4K EEV CCD. The size of the CCD pixels is 15 μm×15 μm. The spatial scale is 0′′.52/pixel in ARGUS direct mode which was used during our observations. The ARGUS array is a rectangular array of 22 by 14 microlenses which is fixed at the center of one positioner arm. We used a spatial scale with a sampling of 0′′.52 per microlens and a total aperture of 11′′.4×7′′.3. The major axis of the array was directed along the major axis of SBS 0335–052E, which is operated by the Association of Universities for Research in Astronomy (AURA) under cooperative agreement with the National Science Foundation (NSF).

The flux calibration has been done using IRAF. The flux-calibrated and redshift-corrected spectrum of the brightest rectangular region delineated by the thick solid line in the center of Fig. 1 is shown in Fig. 2. The region includes clusters 1 and 2. Many strong permitted and forbidden emission lines are seen in the spectrum of this region. One important feature of this spectrum is that its spectral resolution is the highest among all other spectra of SBS 0335–052E obtained so far. In particular, several blends are resolved for the first time for this object, most notably, [O II] λ3726 and [O II] λ3729, He I λ3965, [Ne III] λ3967 and H7 λ3970, [Ar IV] λ4711 and He I λ4713. The spectrum contains emission lines of ions of a broad range of ionisation stages. In particular, [Fe II], [Fe III], [Fe IV], [Fe V], [Fe VI] and [Fe VII] emission lines are detected.

The spectrum of another region centered on the clusters 4 and 5 and delineated by the thick dashed line in Fig. 1 is shown in Fig. 3. The level of continuum in this spectrum is comparable to that in the spectrum of the brightest region (Fig. 2), but emission lines are weaker.

The emission line fluxes and widths were measured in each of the 22×14 lens array using the routine SPLIT in IRAF. Flux errors were derived from the photon statistics using non-flux-calibrated spectra. These errors were propagated in the determination of the electron temperatures, electron number densities and element abundances.

3. Morphology in Continuum and Emission Lines

One of the advantages of the SBS 0335–052E panoramic observations with GIRAFFE/ARGUS is that the spectra of each region with an angular size of 0′′.52×0′′.52 within an aperture 11′′.4×7′′.3 were obtained. This allows us to study the morphology of the galaxy in the continuum and individual emission lines and to construct the model of its H II region.

| Date       | Setup | Wavelength range, Å | Resolving power R | Exposure, s | Airmass | Seeing, arcsec |
|------------|-------|---------------------|-------------------|-------------|---------|----------------|
| 19 Nov 2003 | LR1   | 3620–4081           | 12800             | 2×1200      | 1.066   | 1.03, 0.9      |
| 20 Nov 2003 | LR2   | 3964–4567           | 10200             | 2×1200      | 1.082   | 1.062, 0.8     |
| 20 Nov 2003 | LR3   | 4501–5078           | 12000             | 2×1350      | 1.061   | 1.080, 1.0     |
| 21 Nov 2003 | LR4   | 5010–5831           | 9600              | 2×1200      | 1.295   | 1.170, 0.8     |
| 21 Nov 2003 | LR5   | 5741–6524           | 11800             | 2×1500      | 1.070   | 1.066, 1.3     |
| 21 Nov 2003 | LR6   | 6438–7184           | 13700             | 2×1200      | 1.067   | 1.103, 1.0     |
| 21 Nov 2003 | LR7   | 7102–8343           | 8900              | 2×1200      | 1.286   | 1.486, 0.8     |
| 22 Nov 2003 | LR8   | 8206–9400           | 10400             | 2×1200      | 1.076   | 1.061, 0.7     |

*The first value is at start of exposure, the second value is at end of exposure.
Fig. 2. Spectrum of the brightest part of SBS 0335–052E shown in Fig. 4 as a rectangular region delineated by thick solid line.
Fig. 3. Spectrum of a region centered on clusters 4 and 5 and shown in Fig. 4 as a square region delineated by thick dashed line.
Fig. 4. Images in the continuum near Hβ emission line (a) and in emission lines [O II] λ3727Å (b), [O III] λ4363Å (c), He II λ4686Å (d), [Ar IV] λ4740Å (e), Hβ λ4861Å (f), [O III] λ5007Å (g), He i λ5876Å (h) and Hα λ6563Å (i). All images are shown in the logarithmic flux scale. Positions of stellar clusters are labelled in panel (a). White cross in each panel shows the location of the region with maximum flux of the Hα λ6563Å emission line.

Fig. 5. Distribution of the Hβ λ4861Å equivalent width (a) and of the Hα λ6563Å equivalent width (b) in Å, with errors indicated in parentheses. Each square region corresponds to 0′.52×0′.52. Only regions in which both EW(Hβ) and EW(Hα) could be measured are shown. Dark squares indicate the locations of clusters 1+2 with the maximum flux of the Hα emission line, grey squares the locations of clusters 4+5.

The central part of SBS 0335–052E containing the brightest clusters has an angular size ≲ 2″ (Fig. 1), which is only ≲ 4 times larger than the angular size of 0′.52 of each ARGUS lens. Therefore, for better viewing we rebinned the images, splitting each pixel in 9 from 0′.52 pixel sizes to 0′.17 pixel sizes linearly interpolating flux values in adjacent 0′.52 pixels. In Fig. 4 are shown the rebinned images of SBS 0335–052E in the continuum near Hβ λ4861Å (a), and in the emission lines [O II] λ3727Å (b), [O III] λ4363Å (c), He II λ4686Å (d), [Ar IV] λ4740Å (e), Hβ λ4861Å (f), [O III] λ5007Å (g), He i λ5876Å (h) and Hα λ6563Å (i). In all panels white crosses denote the pixel with the largest flux of the Hα λ6563Å emission line which is coincident with the location of clusters 1+2 in Fig. 1.

The image in the continuum (Fig. 1a) with labelled clusters resembles well the HST UV image (Fig. 1a) despite
the much lower angular resolution determined by a seeing of \(\sim 1''\) (Table 4). Several clusters are seen. However, the angular resolution in the GIRAFFE data is not enough to separate clusters 1 and 2, and 4 and 5.

The images in all emission lines (except for the He II \(\lambda 4686\) line) are very similar. They show very bright emission in the region of clusters 1, 2 and much fainter emission in the direction of other clusters. Furthermore, the equivalent widths of H\(\beta\) \(\lambda 4861\) and H\(\alpha\) \(\lambda 6563\) emission lines in the clusters 1 and 2 (dark squares in Fig. 5) and in the regions around these clusters are high. These facts suggest that clusters 1 and 2 are young, with an age 3–4 Myr, and contain numerous hot and massive ionising main-sequence stars. It is likely, that clusters 7 and probably 3 are also young because EW(H\(\beta\)) and EW(H\(\alpha\)) are high. However, the number of ionising massive stars in those clusters is much lower than in clusters 1 and 2. On the other hand, clusters 4, 5 and 6 are probably more evolved as evidenced by their lower equivalent widths of H\(\beta\) \(\lambda 4861\) and H\(\alpha\) \(\lambda 6563\) emission lines. In particular, the H\(\beta\) equivalent width EW(H\(\beta\)) = 46\(\AA\) for clusters 4, 5 (grey square in Fig. 5) and 87\(\AA\) for the square region delineated by the thick dashed line correspond to an age of \(\sim 6 - 8\) Myr and \(\sim 5\) Myr adopting heavy element mass fraction \(Z = 0.004\) corresponding to the metallicity of the ionised gas is adopted. The larger age of clusters 4, 5 is larger if lower \(Z = 0.0004\) corresponding to the metallicity of the ionised gas is adopted. The young age of clusters 4 and 5 is supported by their weak Po emission as compared to that of clusters 1 and 2 (Thomson et al. 2003). The different age of clusters 1+2 and 4+5 could in principle explain why the brightness of clusters 4 and 5 in the UV range is greater than that of clusters 1 and 2 (Fig. 5), why their brightness is comparable in the optical continuum (Fig. 1) and why clusters 4 and 5 are fainter in the NIR (Thomson et al. 2003). This is because the relative contribution of the ionised gas emission is increased from the UV to the NIR. The effect is stronger for clusters 1 and 2 because of the higher EW(H\(\beta\)) and hence of higher contribution of the ionised gas to the total emission. In addition, the interstellar extinction may play role if it is higher for clusters 1 and 2.

The morphology of SBS 0335–052E in the He II \(\lambda 4686\) emission line (Fig. 5) is significantly different from that in other emission lines. The emission of this line in the direction of the clusters 3 and 4+5 is stronger than that in the direction of the clusters 1 and 2. This offset of the He II \(\lambda 4686\) emission line relative to other nebular lines was noted earlier by Izotov et al. (1999b) and Izotov et al. (2001b). Thus, it is evident that the hard ionising radiation responsible for the He II \(\lambda 4686\) emission is not connected with the young main-sequence stars, but rather related to the post-main-sequence stars or their remnants. This effect is more clearly seen in Fig. 6 where we show the distribution of the relative flux He II \(\lambda 4686/H\beta\). In the direction on the clusters 1 and 2 the relative He II \(\lambda 4686/H\beta\) flux is \(\sim 1–2\%\) while in north-west regions it increases to \(\sim 6–7\%\). Such high relative fluxes of the He II \(\lambda 4686\) emission line is difficult to explain by the ionising stellar radiation (Izotov et al. 2004; Thuan & Izotov 2003, e.g.). Although a small number of Wolf–Rayet stars are found in cluster 3 (see Panaderos et al. 2002) and this paper), other mechanisms such as radiative shocks need to be invoked.

### 4. Kinematics of the Ionised Gas

The second advantage of the present GIRAFFE spectra for SBS 0335–052E is that they are obtained with sufficient enough spectral resolution. Therefore, the panoramic spectroscopic data can be used to study the kinematics of the H II regions in this galaxy.

In Fig. 7 we show the profiles of the H\(\alpha\) \(\lambda 6563\) (a), He II \(\lambda 4686\) (b), [O II] \(\lambda\lambda 3726, 3729\) (c) and [O III] \(\lambda 4363\) (d) emission lines in each pixel of the ARGUS array. Dotted lines show the wavelengths of emission lines adopting the average redshift derived from the observed wavelengths of all strong emission lines in the spectrum of the brightest rectangular region delineated by a thick solid line in Fig. 4. H\(\alpha\) \(\lambda 6563\) is the strongest line in all ARGUS array spectra and thus allows to study the kinematical properties in the low-intensity extended regions of SBS 0335–052E. Other lines originate in different zones of the H II region: He II \(\lambda 4686\) is a characteristic of the highest ionisation zone, [O II] \(\lambda\lambda 3726, 3729\) are characteristics of the lowest ionisation zone, and [O III] \(\lambda 4363\) emission corresponds to the intermediate zone which is overlapped with the highest- and lowest-ionisation zones.

H\(\alpha\) emission is seen almost in the whole region observed with GIRAFFE (Fig. 4a). The total aperture \(11''\times 7.3''\) of ARGUS corresponds to linear size \(\sim 3.1\) kpc\(\times 1.8\) kpc adopting the distance to SBS 0335–052E of 54.3 Mpc (Izotov et al. 1999b). Thus, the observed region is only a part of much larger H II region with a size of \(\sim 6–8\) kpc detected by Izotov et al. (2001b). In the brightest central region and in the slice oriented west-east the H\(\alpha\) \(\lambda 6563\) line is narrow and no systematic offset of the line profile from the dotted line is seen. Thus, no
evidence is present for the rotation of the observed part of the galaxy, since the west-east orientation of the region with narrow profiles is close to that for the disk-like H i cloud seen edge-on (Pustilnik et al., 2001). On the other hand, in the north-south direction the Hα profiles are much broader and with more complex structure, suggesting an outflow of ionised gas in the direction perpendicular to the H i disk. The schematic Hα kinematic model is shown in Fig. 8 where the grey rectangular region is the region with narrow Hα profiles which is oriented approximately along the H i cloud detected by Pustilnik et al. (2001). Two regions with the ionised gas outflow are shown to the north and south of the region with the narrow Hα line. The double-peaked Hα profiles in the northern and southern parts of Fig. 7 suggest the presence of expanding shells of ionised hydrogen with radial components of velocities of ∼ 50 km s\(^{-1}\). This finding is consistent with the presence of a shell of ionised gas at an angular distance of ∼ 5° to the north from the cluster 1 (Thuan et al., 1993). The width of Hα in clusters 4, 5 (grey square in Fig. 7a) is ∼ 2 times larger than that in clusters 1 and 2 (dark square in Fig. 7a) implying higher dynamic activity of the interstellar medium in the former, older clusters. This higher activity is probably due to supernova explosions.

All observed profiles were fit with a single gaussian profile. In Fig. 9a we show FWHMs of the Hα emission line in km s\(^{-1}\) due to the macroscopic motion only. Errors are given in parentheses. The widths are obtained after correction for the instrumental profiles and for the thermal motion, following the formula

\[
\text{FWHM}_{\text{sur}}^2 = \text{FWHM}_{\text{obs}}^2 - \text{FWHM}_{\text{th}}^2 - \text{FWHM}_{\text{th}}^2,
\]

where FWHM\(_{\text{sur}}\), FWHM\(_{\text{obs}}\), FWHM\(_{\text{th}}\) are the width of the macroscopic motion, the observed width, the width of the instrumental profile and the width of the thermal motion, respectively. The FWHM\(_{\text{obs}}\) is obtained from the profiles of night sky emission lines and it is ∼ 25 km s\(^{-1}\). The width of the thermal profile of 30 km s\(^{-1}\) is derived for the electron temperature \(T_\text{e} = 20000\)K and it is applied to the Hα λ4863Å, He II λ4686Å and [O III] λλ4959,5007Å emission lines, while the width of 26 km s\(^{-1}\) is obtained for \(T_\text{e} = 15000\)K and it is applied to the [O II] λλ3729,3726Å emission line.

In the brightest part of SBS 0335–052E the FWHM(Hα) in and around the clusters 1 and 2 is ∼ 40 km s\(^{-1}\). This value is lower than the FWHM(Hα) of 83 km s\(^{-1}\) and 62 km s\(^{-1}\) derived by Petroesian et al. (1997) for the NW and SE components of I Zw 18. Apparently, the dispersion of macroscopic motion is strongly dependent on the age of starburst and it is larger in the evolved starbursts where the SNe activity is higher. Indeed, the difference between clusters 1 and 2 in SBS 0335–052E and the NW and SE components of I Zw 18 is that the clusters 1 and 2 are younger as the equivalent width of the Hβ emission line there, EW(Hβ) ∼ 300–400Å (Fig. 5c), is much larger than the EW(Hβ) ≤ 100Å in the NW and SE components of I Zw 18. It is probable that SNe have not yet appeared or their number is small in clusters 1 and 2.

On the other hand, the macroscopic motion in and around the older clusters 4 and 5 is significantly larger with FWHM(Hα) ∼ 100 km s\(^{-1}\), likely due to a high SN activity. Similar FWHM values and a similar spatial behaviour are found for the [O II] λλ3729,5007Å and [O III] λλ4959,5007Å emission lines. (Figs. 9c and 9d). The tendency of higher FWHM in more evolved clusters is also retained for the He II λ4686Å emission line (Fig. 9e). However, the macroscopic velocity in the regions of He II emission is significantly larger than that in the emission regions of other lines. This difference may be an additional indication that the source of hard radiation is connected with fast radiative shocks.

5. Physical Conditions and Element Abundances

The large aperture (11′′×7′′3), high enough spectral resolution and large wavelength coverage of the ARGUS observations allow the detailed study of physical conditions (electron temperature and electron number density) and element abundances in the H II region. However, there are some limitations to this study. First, observations in different wavelength ranges have been done not in single but in separate exposures during several nights (see Table I). Therefore, due to the varying weather conditions (seeing, transparency), effects of the atmospheric refraction and non-perfect pointing during different exposures the spectra in small apertures such as in a single pixel of 0′′52×0′′52 and in different wavelength ranges are not quite well adjusted since they represent slightly different regions of the galaxy. These effects tend to be lower with increasing aperture. Therefore, depending on the adopted aperture, we will follow different approaches in the determination of element abundances. We consider the element abundance determination from the spectra obtained in apertures with three different sizes. First, for the spectra obtained within the largest available aperture of 11′′×7′′3 there is no need to adjust the different wavelength ranges. This allows to correct consistently the spectra for interstellar extinction using the observed decrement of hydrogen Balmer lines and then derive element abundances. Second, we consider the spectrum of the brightest region of SBS 0335–052E obtained within an aperture 1′′56×1′′04 (the rectangular region delineated by a thick solid line in Fig. 7c). This spectrum is shown in Fig. 9. Since the aperture for this spectrum is relatively small, some adjustment of adjacent wavelength ranges is needed. However, thanks to the high brightness of this region many bright emission lines are seen in its spectrum. Therefore we use the same brightest emission lines in the overlapping wavelength ranges (where this is possible) to scale the spectra in the adjacent wavelength ranges. These lines are Hγ λ3797Å, H δ λ4471Å, [O III] λλ4959,5007Å, [O I] λλ6300,6363Å and He I λλ7065,7067Å. In the remaining two overlapping wavelength ranges λλ5650–5750 and λλ8100–8200 where no strong emission lines are seen we used the continuum levels to adjust
Fig. 7. Profiles of the emission lines: (a) Hα λ6563Å, (b) Heii λ4686Å, (c) [O ii] λ3726,3729Å and (d) [O iii] λ4363Å. All 0′′.52×0′′.52 regions are shown. Dark squares mark the location of clusters 1+2, grey squares the location of clusters 4+5. Thick solid and thick dashed lines in (a) delineate respectively the brightest rectangular region and the second brightest square region for which the spectra are shown in Fig. 2 and 3.

the adjacent spectra. Thus, in the spectrum of the brightest region the determination of the interstellar extinction is still possible, which was used to correct the spectra. Third, in the case of smallest apertures of 0′′.52×0′′.52, in general the signal-to-noise of the spectra is not enough to use the same emission lines in the overlapping wavelength ranges. Therefore, for these apertures we adjust spectra in different wavelength ranges assuming that the ratios of hydrogen Balmer lines correspond to the theoretical values at the electron temperature of Te = 20000K. Hence, no determination of the interstellar extinction is possible for the smallest apertures and not all wavelength ranges could be adjusted. Fortunately, it is possible to adjust wavelength ranges containing the [O ii] λ3726, 3729Å, [O iii] λ4363, 4959, 5007Å, [S ii] λ6717, 6731Å emission lines, therefore at least the determination of the electron temperature Te(O iii), the electron number density Ne(S ii) and the oxygen abundance is possible.

To derive Te, Ne and heavy element abundances we follow the prescriptions by Izotov et al. (2006a). Namely, where possible, the coefficient of interstellar extinction C(Hβ) and the equivalent width of absorption hydrogen lines EWabs are derived from the observed hydrogen Balmer decrement. In this procedure we assume that EWabs is the same for all hydrogen lines. Then the fluxes of emission lines were corrected for interstellar extinction and underlying stellar absorption (where this is possible).

We adopt the three zone model of the H ii region. The electron temperature Te(O iii) in the high-ionisation zone is derived from the [O iii] λ4363/(λ4959+λ5007) flux ratio. This temperature is used to derive abundances of ions O2+ and Ne2+. Since Heii λ4686Å emission is present in the SBS 0335–052E spectrum, the O3+ abundance is derived following Izotov et al. (2006a) and adopting Te(O iii). Since the O3+ abundance is significantly lower than the O2+ abundance, the uncertainties in the temperature for the zone where the O3+ ion is present introduce only a small uncertainty in the total oxygen abundance. Some other emission lines of high-ionisation ions Ar3+, Cl3+, Fe3+–Fe6+ are seen in the spectrum of SBS 0335–052E in Fig. 2. In general these ions are present in the inner part of the H ii region with a temperature higher than Te(O iii).

However, since there is no temperature constraint from observations for these ions and atomic data are not well...
known for some of them, we decided not to use these ions for the abundance determination. For the intermediate-ionisation zone we adopt the electron temperature for the abundance determination. For the intermediate-known for some of them, we decided not to use these ions and ionised gas and their abundances in the H are subject to large uncertainties.

We first consider the chemical composition in the brightest-5.1. Heavy Element Abundances in the Brightest Region

We first consider the chemical composition in the bright- in Table 4. The extinction coefficient $C$(Hβ), the equiva- lent width $EW_{abs}$ of the absorption hydrogen Balmer lines and the observed flux $F$(Hβ) of the Hβ emission line are shown at the end of Table 4. Electron temperatures, electron number densities, ionic and total heavy element abundances for the brightest region are shown in the second column of Table 2. In general, the derived parameters are consistent with previous determinations e.g. by Izotov et al. (1997b), Izotov et al. (1999), Thuan & Izotov (2007). In particular, the electron temperature $T_e$(O iii) in all measurements is high and is close to 20000K. It was found in previous studies that the H ii region in SBS 0335–052E is relatively dense. We confirm this finding. The electron number density, which we derive from the [S ii] $\lambda 6717/\lambda 6731$ flux ratio in the brightest region, is $\sim 150$ cm$^{-3}$. A similar value is obtained from the [O ii] $\lambda 3726/\lambda 3729$ flux ratio. On the other hand, the electron number density derived from the [Ar iv] $\lambda 4711/\lambda 4740$ flux ratio, using the data from Aller (1984), is much larger, $\sim 6000$ cm$^{-3}$. Thus, it appears that the high-ionisation regions are much denser than the low-ionisation regions.

The oxygen abundance $12+\log$ O/H = 7.31 ± 0.01 is in perfect agreement with recent determinations by Izotov et al. (1997b), Izotov et al. (1999) and Thuan & Izotov (2007). The Ne/O, S/O, Cl/O and Ar/O abundance ratios are very close to the average values found by, e.g., Izotov et al. (2006a) for the large sample of low-metallicity emission-line galaxies. On the other hand, the N/O abundance ratio $\sim 1.37$ appears higher than the mean value of $\sim 1.5$ to $\sim 1.6$ for the most metal-deficient BCDs (Izotov & Thuan, 1999; Izotov et al. 2006a). Since only N$^+$ lines are observed in the optical spectrum of SBS 0335–052E the total nitrogen abundance is derived as N/H = $ICF$(N)$\times$N$^+$/H$^+$, where $ICF$(N) $\sim$
(O$^{3+}$+O$^{2+}$+O$^+$/O$^+$. Inspection of Table 4 shows that the relative flux [O ii] $\lambda$3727/H$\beta$ is 0.2, or 30% lower than that in some other observations of SBS 0335–052E (e.g. Izotov et al., 1999; Papaderos et al., 2006) resulting in high ICF(N). The lower [O ii] $\lambda$3727 flux in the bright region (Table 4) could be due to observational uncertainties (slightly different pointings of SBS 0335–052E during observations with different setups, effect of differential refraction, variable seeing, etc.). Adopting a [O ii] $\lambda$3727 flux $\sim$30% higher will result in log N/O $\sim$1.5, in much better agreement with other determinations. Such increase of the [O ii] $\lambda$3727 flux will also slightly decrease by $\sim$0.1 dex the iron abundance, while the abundances of other heavy elements will remain almost unchanged. In particular, the oxygen abundance 12 + log O/H will be increased only by 0.01 dex. The Fe/O abundance ratio is high and is typical for the extremely metal-deficient BCDs (Izotov et al., 2006a). This fact suggests that the depletion of iron onto dust in SBS 0335–052E is small.

### 5.2. Heavy Element Abundances from the Integrated Spectrum of SBS 0335–052E

We use panoramic VLT/GIRAFFE data to obtain the integrated spectrum of SBS 0335–052E by summing every of 22×14 spectra in the whole aperture. The resulting spectrum is significantly more noisy as compared to the spectrum of the brightest region because many spectra of low-brightness regions were co-added to the spectrum of the brightest region. However, the integrated spectrum is not subject to the observational uncertainties which are much more important for the spectra obtained with the smaller apertures (non-perfect pointing, variable seeing). Additionally, it allows to obtain integrated characteristics such as the luminosity of the galaxy in individual lines.

In Table 5 are shown the measured absolute fluxes $F(\lambda)$ of the emission lines, the absolute fluxes $F(\lambda)$ corrected for the interstellar extinction and underlying stellar absorption, the respective fluxes relative to the H$\beta$ $\lambda$4861 flux, $F(\lambda)/F(H\beta)$ and $I(\lambda)/I(H\beta)$, the equivalent widths $EW(\lambda)$ of the emission lines, the interstellar extinction coefficient $C(H\beta)$ and equivalent width of hydrogen Balmer absorption lines EW$_{abs}$. The absolute measured flux of H$\alpha$ emission line $F(H\alpha) = 3.45 \times 10^{-13}$ erg s$^{-1}$ cm$^{-2}$ is consistent with the value $3.23 \times 10^{-13}$ erg s$^{-1}$ cm$^{-2}$ obtained by Pustilnik et al. (2004). The luminosities of the H$\beta$ and H$\alpha$ emission lines corrected for interstellar extinction and underlying stellar absorption are equal to $L(H\beta) = 6.4 \times 10^{40}$ erg s$^{-1}$ and $L(H\alpha) = 1.8 \times 10^{41}$ erg s$^{-1}$, corresponding to the equivalent number of O7V stars $N(O7V) = 1.3 \times 10^4$. The major fraction of these stars ($\gtrsim 90\%$) is located in the two compact clusters 1 and 2 as it is evidenced by the high-resolution spatial distribution of the Pa$\alpha$ emission in SBS 0335–052E obtained by Thompson et al. (2006) from the HST observations. To the best of our knowledge, these two clusters (most likely, cluster 1) are among the richest super-star clusters, hosting a very large number of O stars within a region of angular size $\lesssim 0'.1 \sim 0'.2$, corresponding to a linear size $\lesssim 25 \sim 50$ pc.

Table: Heavy Element Abundances

| Property       | Bright | Total |
|----------------|--------|-------|
| $T_e$ (O iii), K | 2020±200 | 20360±480 |
| $T_e$ (O ii), K | 15620±170 | 15630±340 |
| $T_e$ (S iii), K | 18750±200 | 18910±400 |
| $N_e$(O ii), cm$^{-3}$ | 120±50 | 150±80 |
| $N_e$(Ar iv), cm$^{-3}$ | 6000±2000 | ... |
| $N_e$(S ii), cm$^{-3}$ | 146±41 | 213±74 |
| $O^+$/H$^+$, ($\times 10^3$) | 0.164±0.005 | 0.211±0.013 |
| $O^{2+}$/H$^+$, ($\times 10^3$) | 1.838±0.054 | 1.664±0.090 |
| $O^{3+}$/H$^+$, ($\times 10^3$) | 0.032±0.001 | 0.081±0.006 |
| O/H, ($\times 10^3$) | 2.034±0.054 | 1.956±0.091 |
| 12+log O/H | 7.31±0.01 | 7.29±0.02 |
| N$^+$/H$^+$, ($\times 10^7$) | 0.779±0.019 | 0.673±0.033 |
| ICF(N) | 11.25 | 8.53 |
| N/H, ($\times 10^7$) | 8.765±0.235 | 7.541±0.304 |
| log N/O | $-1.37±0.02$ | $-1.53±0.03$ |
| Ne$^{2+}$/H$^+$, ($\times 10^6$) | 2.977±0.092 | 2.941±0.199 |
| ICF(Ne) | 1.04 | 1.06 |
| Ne/H, ($\times 10^6$) | 3.099±0.102 | 3.126±0.234 |
| log Ne/O | $-0.82±0.02$ | $-0.80±0.04$ |
| S$^{+}$/H$^+$, ($\times 10^7$) | 0.314±0.006 | 0.348±0.013 |
| S$^{2+}$/H$^+$, ($\times 10^7$) | 1.539±0.052 | 1.433±0.136 |
| ICF(S) | 2.36 | 1.93 |
| S/H, ($\times 10^7$) | 4.363±0.122 | 3.431±0.264 |
| log S/O | $-1.67±0.02$ | $-1.76±0.04$ |
| Cl$^{+}$/H$^+$, ($\times 10^9$) | 2.846±0.560 | ... |
| ICF(Cl) | 1.48 | ... |
| Cl/H, ($\times 10^9$) | 4.207±0.827 | ... |
| log Cl/O | $-3.68±0.09$ | ... |
| Ar$^{2+}$/H$^+$, ($\times 10^6$) | 4.478±0.094 | 4.408±0.199 |
| ICF(Ar) | 1.49 | 1.34 |
| Ar/H, ($\times 10^6$) | 6.680±0.287 | 5.900±0.812 |
| log Ar/O | $-2.48±0.02$ | $-2.52±0.06$ |
| Fe$^{2+}$/H$^+$, ($\times 10^7$) | 0.453±0.033 | ... |
| ICF(Fe) | 17.02 | ... |
| Fe/H, ($\times 10^7$) | 7.717±0.567 | ... |
| log Fe/O | $-1.42±0.03$ | ... |
for most metal-deficient galaxies (Izotov & Thuan 1999; Izotov et al. 2006a).

5.3. Distribution of the Physical Parameters and Oxygen Abundance

The VLT/GIRAFFE panoramic spectra allow also to study the distribution of the electron temperature $T_e$ (O III), the electron number density $N_e$ (S II) and heavy element abundances in the H II region. For this we use the spectra obtained for the 0′52×0′52 apertures. We took in consideration only spectra in which at least the following lines of heavy elements are detected: [O II] λ3726, 3729 Å, [O III] λ4363, 4959, 5007 Å. This allows to derive the electron temperature $T_e$ (O III) and oxygen abundance. From these spectra we excluded those spectra, where the oxygen abundance 12 + log O/H is derived with an error greater than 0.1 dex.

In Fig. 10, is shown the distribution of the electron temperature $T_e$ (O III). It is seen that the H II region is hot in all small apertures and has the characteristic temperature of ∼20000K. There is a slight spatial trend of the electron temperature with $T_e$ (O III) being slightly higher in the western part and slightly lower in the eastern part. The electron number density derived from the [S II] λ6717, 6731 Å emission lines is high, of several hundred particles per cm$^3$ (Fig. 11). However, the errors in the determination of $N_e$ are large and are caused by the low intensity of the [S II] emission lines. Similar number densities are derived from the [O II] λ3726/λ3729 flux ratio. Although [O II] λ3726, 3729 Å emission lines are brighter than [S II] λ6717, 6731 Å, the low signal-to-noise ratio of the spectrum containing [O II] lines (Fig. 2) due to lower sensitivity of the GIRAFFE detector in that wavelength range prevents the determination of the electron number density from the [O II] lines in a region larger than that from the [S II] emission lines. As already mentioned, other emission lines detected in our spectra can in principle be used to determine the electron density. [Ar IV] λ4711 and 4740 Å are strong enough only in the brightest region of SBS 0335–052E, where they indicate consistently an electron number density in the range 10$^4$ – 10$^5$ cm$^{-3}$. [Cl III] λ5517 and 5537 Å are too weak even in the brightest region (Fig. 2) and are therefore not used.

The oxygen abundance 12 + log O/H distribution is shown in Fig. 11. There are some variations of the oxygen abundance in the range 7.00 – 7.42 with a slight trend of decreasing of 12 + log O/H from the East to the West. In particular, it appears that the oxygen abundance in cluster 7 of ∼7.2 (western side of Fig. 10) is slightly lower than in other clusters, confirming the finding by Papaderos et al. (2000). Thus, there is some evidence for a possible self-enrichment by heavy elements (cf. Izotov et al. 1997b, 1999) or for the presence of “initial” abundance variations in the gas. However, we point out here that the errors in the electron temperature, electron number density and oxygen abundance include only errors derived based on the photon statistics of non-flux-calibrated emission line fluxes and they do not take into account uncertainties in pointing, variable seeing, differential refraction, etc., which is difficult to estimate. Therefore, variations in the oxygen abundance may not be statistically significant.

In Fig. 10a are shown variations of N/O abundance ratio. The log N/O varies in the wide enough range from
−1.13 to −1.65 with relatively small errors. However, the real errors might be much higher because of the limitations introduced by small aperture sizes of 0′.52 × 0′.52 for each spectrum. The same is true for the distribution of the Ne/O abundance ratio (Fig. 10).

### 5.4. Helium Abundance from the Integrated Spectrum

SBS 0335–052E being one of the most metal-deficient BCDs plays an important role in the determination of the primordial He mass fraction \( Y_p \) and, thus, in the determination of the baryonic mass fraction of the Universe. Since the precision in the determination of \( Y_p \) should be better than \( \sim 1\% \) to put useful constraints on the cosmological models, high signal-to-noise spectra are needed for this. Additionally, several systematic effects should be taken into account, and spectra and emission line fluxes should be corrected for them (see for details Izotov & Thuan 2004; Izotov et al. 2006b). These are the corrections for (1) interstellar extinction, (2) ionisation structure, (3) collisional excitation of helium lines, (4) fluorescence in helium lines, (5) temperature variations, (6) underlying stellar absorption, (7) collisional excitation of hydrogen lines. All these corrections are at a level of a few percent and, apart from (2) influence each other in a complicated way. The case of SBS 0335–052E is particularly complicated, because its H II region is dense, hot and optically thick in some He I emission lines (Izotov et al. 1999). Therefore, effects (3), (4) and (7) are strong in the H II region of SBS 0335–052E.

To derive the He abundance we use the integrated spectrum of SBS 0335–052E because it is least dependent on the observational parameters discussed above. The He\(^{+}\) abundance \( y^+ \) which is derived from the He I emission lines depends on the adopted He I emissivities. We adopt the new He I emissivities by Porter et al. (2005). In this paper, following Izotov et al. (1994, 1997a, 2000b) and Izotov et al. (2006b), we use the five strongest He I λ3889, λ4471, λ5876, λ6678 and λ7065 emission lines to derive \( N_e(\text{He}^+\text{)} \) and \( \tau(\lambda3889) \). The He I λ3889 and λ7065 lines play an important role because they are particularly sensitive to both quantities. Since the He I λ3889 line is blended with the H8 λ3889 line, we have subtracted the latter, assuming its intensity to be equal to 0.107 I(Hβ) (Aller 1984).

Besides the emissivities the derived \( y^+\text{He}^+/\text{H}^+ \) abundance depends on several other parameters: collisional excitation of helium emission lines, electron number density \( N_e(\text{He}^+\text{)} \) and electron temperature \( T_e(\text{He}^+\text{)} \), equivalent widths EW(λ3889), EW(λ4471), EW(λ5876), EW(λ6678) and EW(λ7065) of He I stellar absorption lines, optical depth \( \tau(\lambda3889) \) of the He I λ3889 emission line. We use Monte Carlo simulations for the \( y^+ \) determination randomly varying the parameters in their ranges. First, we subtract the fractions of the Hα and Hβ observed fluxes due to the collisional excitation. We adopt that the fraction \( \Delta I_{\text{coll}}(\text{Hα})/I(\text{Hα}) \) of the Hα flux due to the collisional excitation varies in the range 0%–5% of the total flux. This fraction is randomly generated 100 times in the adopted range. The fraction of the Hβ emission line flux due to the collisional excitation is adopted to be three times less than that of the Hα flux. For each generated fraction of the Hα and Hβ they are subtracted from the total observed fluxes and then all emission line fluxes are corrected for the interstellar extinction and abundances of elements are calculated.

To calculate \( y^+ \) we simultaneously and randomly vary \( N_e(\text{He}^+\text{)} \), \( T_e(\text{He}^+\text{)} \) and \( \tau(\lambda3889) \). The total number of such realizations is \( 10^5 \) for each value of \( \Delta I_{\text{coll}}(\text{Hα})/I(\text{Hα}) \). Thus, the total number of Monte Carlo realizations is \( 100 \times 10^5 = 10^7 \). As for He I underlying stellar absorption, we keep constant values of EW(λ3889), EW(λ4471), EW(λ5876), EW(λ6678) and EW(λ7065) during simulations. In our spectrum other He I emission lines, namely He I λ3820, λ4388, λ4026, λ4921, and λ7281 are seen. However, we do not attempt to use these lines for He abundance determination because they are much weaker as compared to the five brightest lines, and hence have larger uncertainties.

We solve the problem by minimization of the expression

\[
\chi^2 = \frac{1}{\sigma^2} \sum \left( \frac{y_i^+ - y_{\text{mean}}^+}{\sigma(y_i^+)} \right)^2 ,
\]

where \( y_i^+ \) is the He\(^{+}\) abundance derived from the flux of the He I emission line with label \( i \), \( \sigma(y_i^+) \) is the statistical error of the He abundance. The quantity \( y_{\text{mean}}^+ \) is the weighted mean of the He\(^{+}\) abundance as derived from the equation

\[
y_{\text{mean}}^+ = \frac{\sum_i k_i y_i^+/\sigma^2(y_i^+)}{\sum_i k_i / \sigma^2(y_i^+)} .
\]

We use all five He I emission lines to calculate \( \chi^2 \) (i.e., \( n = 5 \)), but only three lines, He I λ4471, He I λ5876 and He I λ6678 to calculate \( y_{\text{mean}}^+ \) (\( k = 3 \)). This is because the fluxes of He I λ3889 and He I λ7065 emission lines are more uncertain compared to other three He I emission lines.

The best solution for \( y_{\text{mean}}^+ \) is found from the minimum of \( \chi^2 \), the systematic error \( \sigma_{\text{sys}} \) is obtained from the dispersion of \( y_{\text{mean}}^+ \) in the range of \( \chi^2_{\text{min}} \) and \( \chi^2_{\text{min}} + 1 \). Then the total error for \( y_{\text{mean}}^+ \) is derived from \( \sigma_{\text{tot}}^2 = \sigma_{\text{stat}}^2 + \sigma_{\text{sys}}^2 \).

Additionally, since the nebular He II λ4686Å emission line was detected, we have added the abundance of doubly ionised helium \( y^{2+} \equiv He^{2+/H^+} \) to \( y^+ \). Although the He\(^{2+}\) zone is hotter than the He\(^+\) zone, we adopt \( T_e(He^{2+/H^+}) = T_e(He^+) \). The last assumption introduces a little change in \( y \) value, because the value of \( y^{2+} \) is small (\( \sim 4\% \) of \( y^+ \)). Finally, the ionisation correction factor \( ICF(\text{He}) \) is taken into account from Izotov et al. (2006b) to convert \( y^+ + y^{2+} \) to the total He abundance \( y = He/H \).

The derived parameters and He abundances for every He I emission line are shown in Table 8 Here,
collisional excitation. It is seen from this Table, that the fraction consistent within the errors) than the primordial He mass
in smaller BCDs (see e.g. Izotov & Thuan, 2004) implying in other BCDs (see e.g. Izotov & Thuan, 2004) implying
important contribution of the fluorescent enhancement of 

| Parameter | Value |
|-----------|-------|
| $T_e$(O III), K | 20180 |
| $T_e$(He$^+$), K | 20010 |
| $N_e$(He$^+$), cm$^{-3}$ | 295 |
| EW$_{abs}$(λ4471), Å | 0.4 |
| EW$_{abs}$(λ3889)/EW$_{abs}$(λ4471) | 1.0 |
| EW$_{abs}$(λ5876)/EW$_{abs}$(λ4471) | 0.3 |
| EW$_{abs}$(λ6678)/EW$_{abs}$(λ4471) | 0.1 |
| EW$_{abs}$(λ7065)/EW$_{abs}$(λ4471) | 0.1 |
| $\tau$(λ3889) | 2.9 |
| $\Delta I_{col}$(Hα)/I(Hα), % | 4.95 |
| $\chi^2_{min}$ | 0.1825 |
| $y^+(\lambda 3889)$ | 0.0782±0.0134 |
| $y^+(\lambda 4471)$ | 0.0802±0.0032 |
| $y^+(\lambda 5876)$ | 0.0788±0.0012 |
| $y^+(\lambda 6678)$ | 0.0787±0.0018 |
| $y^+(\lambda 7065)$ | 0.0789±0.0018 |
| $y^+(\text{weighted mean})$ | 0.2463±0.0030 |

The derived weighted mean He mass fraction in SBS 0335–052E, $Y = 0.2463 ± 0.0030$, is slightly lower (but consistent within the errors) than the primordial He mass fraction $Y_p = 0.24815 ± 0.0003 ± 0.0006$ (syst.) from the 3-year data of the WMAP experiment (Spergel et al., 2006) and the D abundance, supporting the standard cosmological model of the primordial nucleosynthesis.

### 6. Wolf-Rayet Stars

The search for Wolf-Rayet (WR) stars in extremely low-metallicity dwarf galaxies is of great interest for constraining stellar evolution models. However, such studies are difficult since the strength of WR emission lines is significantly reduced with decreasing metallicity. Therefore, very high signal-to-noise ratio spectra are required to detect weak WR features. For a long time no WR galaxies with an oxygen abundance $12 + \log O/H < 7.9$ were known

In principle a more detailed search for WR stars in SBS 0335–052E, a more precise localisation in the galaxy, and also to resolve the λ4650 WR bump into N III λ4640 and C IV λ4658 broad features, thus allowing the detection of both late nitrogen WR stars (WNL stars) and early carbon WR stars (WCE stars). Previous observations by Papaderos et al. (2006) had too low spectral resolution to definitely distinguish between two types of WR stars. However, there are some limitations of GIRAFFE/ARGUS observations which make such study more difficult as compared with that of Papaderos et al. (2006).

GIRAFFE/ARGUS observations allow in principle a more detailed search for WR stars in SBS 0335–052E, a more precise localisation in the galaxy, and also to resolve the λ4650 WR bump into N III λ4640 and C IV λ4658 broad features, thus allowing the detection of both late nitrogen WR stars (WNL stars) and early carbon WR stars (WCE stars). Previous observations by Papaderos et al. (2006) had too low spectral resolution to definitely distinguish between two types of WR stars. However, there are some limitations of GIRAFFE/ARGUS observations which make such study more difficult as compared with that of Papaderos et al. (2006). Although Papaderos et al. (2006) have observed with the smaller 3.6m ESO telescope, their spectrum has a higher S/N ratio because of the ~3 times longer exposure and ~10 times lower spectral resolution.

We checked all 0′52×0′52 spectra obtained with GIRAFFE and found that broad WR features near λ4650 are likely present only in one spectrum associated with cluster 3. This spectrum is shown in Fig. 11. The S/N ratio of ~5 of this spectrum is not high, but broad WR features are clearly seen. Thus, we confirm the finding of Masegosa et al. (1991). Later, Izotov et al. (1997c) and Legrand et al. (1997) have discovered WR stars in I Zw 18, at that time the most metal-deficient emission-line galaxy with the oxygen abundance $12 + \log O/H = 7.17$. Thus, it appears that WR stars could be found in any other dwarf emission-line galaxy with active star formation, including SBS 0335–052E. However, the strength of WR emission features depends not only on metallicity but also on the age of a starburst. In starbursts with the metallicity of SBS 0335–052E, the WR stage is expected to be very short, typically less than 1 Myr (Schaerer & Vacca, 1998; de Mello et al., 1998). Therefore, not all young clusters ionising the interstellar medium in SBS 0335–052E may be expected to contain WR stars. Recently, Papaderos et al. (2006) have found that WR stars of the early carbon sequence (WC4 stars) are present in cluster 3 of SBS 0335–052E.

Fig. 11. Part of the spectrum of cluster 3 showing the probable broad Wolf-Rayet emission lines N III λ4640 and C IV λ4658 labelled "WR".
by Papaderos et al. (2000) that WR stars appear to be present in cluster 3. However, we find that the WR feature consists in fact of two lines: N III λ4640 and C IV λ4658. The latter line is blended with the much narrower nebular [Fe III] λ4658 emission line. We find that, after subtraction of the [Fe III] line from the λ4658 blend, the fluxes of the N III λ4640 and C IV λ4658 lines and their FWHMs are similar, \( \sim (4\pm1) \times 10^{-17} \) erg s\(^{-1}\) cm\(^{-2}\) and 6.5Å, respectively. Thus, the total flux of the λ4650 bump (N III λ4640 + C IV λ4658) non-corrected for extinction is \( \sim 8 \times 10^{-17} \) erg s\(^{-1}\) cm\(^{-2}\), or 2/3 that measured by Papaderos et al. (2006) in a larger aperture. The total equivalent width of this bump is EW(λ4650) \( \sim 9 \)Å. The C IV λ4658 emission line should be accompanied by the C IV λ5808 emission line with a comparable flux. Unfortunately, the redshifted C IV λ5808 emission line in SBS 0335–052E coincides with the night sky Na I λ5890, 5895 emission lines. Therefore, the imperfect night sky subtraction hinders the detection of the WR line.

The observed (i.e. not extinction corrected) N III λ4640 and C IV λ4658 emission line luminosities of cluster 3 are equal to \( L(\text{N III} \lambda4640) = L(\text{C IV} \lambda4658) = 1.4 \times 10^{37} \) erg s\(^{-1}\) and correspond to the number of WNL and stars N(WNL) \( \sim 35 \pm 8 \) and to the number of WC4 stars \( N(\text{WC4}) \sim 3 \pm 1 \) if we adopt the “standard” WR line luminosities computed by Schaerer & Vacca (1998), i.e. assuming N III λ4640 and C IV λ4658 emission line luminosities of 4\( \times 10^{35} \) erg s\(^{-1}\) and 5\( \times 10^{36} \) erg s\(^{-1}\) respectively for a single WNL and WC4 star. The derived values of WNL and WC4 stars are to be taken with a grain of salt for various reasons. First these are likely lower limits because of the neglected reddening, and since WR stars at low metallicity may have lower intrinsic line luminosities (e.g. Crowther & Hadfield, 2006). Second, it is not necessarily clear that the observed lines are indeed due to WN and WC stars, as their strengths and widths are somewhat unusual. Their relatively small widths could be due to lower mass loss rates and/or smaller wind velocities in WR stars at such low metallicity (Crowther et al., 2002; Vink & de Koter, 2003; Crowther & Hadfield, 2006). Alternatively, from their relative strength and the small widths, the observed lines could also correspond to very late WNL stars (WN 10h-11h) (cf. Crowther & Smith, 1997), if most of the C IV λ4658 was nebular. Given the faintness of these spectral signatures and the lack of known individual WR stars at such low metallicities as comparison objects, it is difficult to draw firm conclusions on the WR content in this cluster.

7. Conclusions

In this paper we present panoramic spectroscopic observations with the VLT/GIRAFFE in the spectral range λ3620–9400Å of one of the most metal-deficient blue compact dwarf (BCD) galaxies, SBS 0335–052E. Our findings can be summarized as follows:

1. The morphology of the galaxy in different lines is very similar, except for the He II λ4686Å line, suggesting that the main ionising source in this galaxy is the compact super-star clusters (SSC) 1+2 in the SE part of the galaxy, most likely the cluster 1. The equivalent number of O7v ionising stars in this cluster is \( \sim 13000 \), making it the most powerful SSC known so far.

2. The maximum emission of the He II λ4686Å line is offset to the older compact clusters 4 and 5. Additionally, its width is 1.5 – 2 times greater than the width of other lines. Furthermore, the width of this and other emission lines is much higher in clusters 4,5 than in cluster 1. These facts imply that the hard ionising radiation responsible for He II emission is most likely produced by fast radiative shocks. However, some contribution from WR stars located in cluster 3 cannot be excluded.

3. The analysis of kinematical properties of the ionised gas in SBS 0335–052E suggests the presence of gas outflow in the direction perpendicular to the galaxy disk.

4. We derive physical parameters in the H II region of SBS 0335–052E, suggesting that the ionised gas is hot (\( \sim 20000K \)) and dense (\( \geq \) several hundred particles per cm\(^{-3}\)). The oxygen abundances 12 + log O/H in the brightest region and in the integrated spectrum of SBS 0335–052E are respectively 7.31±0.01 and 7.29±0.02, and they are consistent with previous determinations. The other heavy element-to-oxygen abundance ratios are consistent with the average values derived for other most metal-deficient galaxies. There is a slight decrease of oxygen abundance from the East to the West suggesting some self-enrichment of the ionised gas by heavy elements.

5. We derive the He mass fraction \( \text{Y} \) from the integrated spectrum taking into account all possible systematic effects in the He abundance determination. Our value \( \text{Y} = 0.2463 \pm 0.0030 \) is slightly lower but consistent within the errors with the primordial He mass fraction determined from the 3 year WMAP data and from primordial deuterium measurements.

6. We confirm the presence of Wolf-Rayet stars in the cluster 3 found previously by Papaderos et al. (2000). The lower limits of the WNL and WC4 star numbers are \( \sim 35 \pm 8 \) and \( \sim 3 \pm 1 \), respectively if standard WR line luminosities are assumed.

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Online Material
| Line         | $100 \times F(\lambda)/F(\text{H}\beta)$ | $100 \times I(\lambda)/I(\text{H}\beta)$ | $E(W(\lambda))$ |
|--------------|------------------------------------------|------------------------------------------|-----------------|
| LR1          |                                          |                                          |                 |
| 3634 He i    | 0.25±0.02                                | 0.29±0.13                                | 0.33±0.04       |
| 3631 H24     | 0.47±0.02                                | 1.08±0.23                                | 0.89±0.06       |
| 3674 H23     | 0.59±0.03                                | 1.21±0.20                                | 1.13±0.07       |
| 3676 H22     | 0.60±0.03                                | 1.21±0.19                                | 1.19±0.07       |
| 3679 H21     | 0.80±0.04                                | 1.36±0.15                                | 1.86±0.08       |
| 3683 H20     | 1.16±0.04                                | 1.76±0.13                                | 2.85±0.10       |
| 3687 H19     | 0.83±0.04                                | 1.33±0.13                                | 2.28±0.10       |
| 3692 H18     | 1.08±0.04                                | 1.62±0.12                                | 2.90±0.09       |
| 3697 H17     | 1.07±0.04                                | 1.70±0.14                                | 3.21±0.10       |
| 3704 H16     | 1.40±0.04                                | 2.06±0.12                                | 3.17±0.09       |
| 3705 He i    | 0.59±0.03                                | 0.68±0.08                                | 1.37±0.08       |
| 3712 H15     | 1.21±0.04                                | 1.83±0.13                                | 2.81±0.09       |
| 3722 H14     | 1.80±0.05                                | 2.53±0.13                                | 3.90±0.10       |
| 3726 [O ii]  | 8.12±0.13                                | 9.26±0.18                                | 17.73±0.14      |
| 3729 [O ii]  | 9.51±0.15                                | 10.85±0.20                               | 20.86±0.17      |
| 3734 H13     | 2.01±0.05                                | 2.76±0.13                                | 4.46±0.13       |
| 3750 H12     | 2.75±0.07                                | 3.59±0.14                                | 6.12±0.15       |
| 3771 H11     | 3.08±0.08                                | 3.96±0.15                                | 6.65±0.14       |
| 3798 H10     | 4.03±0.10                                | 5.04±0.17                                | 8.45±0.19       |
| 3820 He i    | 0.72±0.07                                | 0.81±0.11                                | 1.63±0.17       |
| 3835 H9      | 6.19±0.15                                | 7.37±0.20                                | 15.35±0.29      |
| 3868 [Ne iii]| 21.29±0.36                               | 23.84±0.43                               | 47.27±0.45      |
| 3889 He i+H8 | 14.13±0.29                               | 16.22±0.36                               | 32.49±0.50      |
| 3965 He i    | 0.61±0.23                                | 0.68±0.26                                | 1.60±0.58       |
| 3967 [Ne iii]| 6.51±0.29                                | 7.20±0.33                                | 18.31±0.70      |
| 3970 H7      | 14.02±0.42                               | 15.85±0.50                               | 39.08±0.93      |
| LR2          |                                          |                                          |                 |
| 3965 He i    | 0.64±0.01                                | 0.71±0.08                                | 1.30±0.04       |
| 3967 [Ne iii]| 7.16±0.11                                | 7.91±0.15                                | 14.64±0.08      |
| 3970 H7      | 13.97±0.20                               | 15.92±0.26                               | 28.70±0.10      |
| 4026 He i    | 1.60±0.03                                | 1.75±0.09                                | 3.38±0.05       |
| 4068 [S ii]  | 0.25±0.02                                | 0.28±0.08                                | 0.54±0.04       |
| 4101 Hδ      | 23.99±0.35                               | 26.53±0.42                               | 51.74±0.16      |
| 4121 He i    | 0.24±0.02                                | 0.26±0.08                                | 0.52±0.05       |
| 4144 He i    | 0.41±0.03                                | 0.44±0.08                                | 0.89±0.06       |
| 4169 He i    | 0.10±0.02                                | 0.10±0.08                                | 0.21±0.05       |
| 4227 [Fe v]  | 0.20±0.03                                | 0.21±0.08                                | 0.45±0.08       |
| 4249 [Fe ii] | 0.07±0.02                                | 0.08±0.07                                | 0.17±0.05       |
| 4287 [Fe ii] | 0.16±0.02                                | 0.17±0.07                                | 0.39±0.07       |
| 4340 Hγ      | 46.80±0.68                               | 49.82±0.75                               | 114.10±0.33     |
| 4359 [Fe ii] | 0.17±0.03                                | 0.18±0.06                                | 0.59±0.07       |
| 4363 [O iii] | 10.52±0.17                               | 11.08±0.19                               | 25.39±0.18      |
| 4368 O i     | 0.10±0.03                                | 0.11±0.07                                | 0.25±0.06       |
| 4379 N iii   | 0.10±0.03                                | 0.10±0.07                                | 0.24±0.08       |
| 4388 He i    | 0.44±0.03                                | 0.46±0.07                                | 1.10±0.10       |
| 4414 [Fe ii] | 0.08±0.02                                | 0.08±0.07                                | 0.19±0.07       |
| 4416 [Fe ii] | 0.05±0.02                                | 0.05±0.06                                | 0.13±0.07       |
| 4438 He i    | 0.11±0.03                                | 0.12±0.07                                | 0.29±0.08       |
| 4471 He i    | 3.75±0.08                                | 3.90±0.10                                | 9.83±0.15       |
| LR3          |                                          |                                          |                 |
| 4452 [Fe ii] | 0.10±0.01                                | 0.10±0.06                                | 0.25±0.04       |
| 4471 He i    | 3.75±0.06                                | 3.90±0.09                                | 9.88±0.08       |
| 4541 He ii   | 0.05±0.01                                | 0.05±0.06                                | 0.13±0.04       |
| 4571 Mg i    | 0.09±0.01                                | 0.10±0.06                                | 0.26±0.04       |
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**Table 4.** Continued.

| Line   | $100 \times F(\lambda)/F(H\beta)$ | $100 \times I(\lambda)/I(H\beta)$ | $EW(\lambda)^a$ |
|--------|-----------------------------------|-----------------------------------|-----------------|
| 4658 [Fe III] | 0.26±0.02                         | 0.26±0.06                         | 0.75±0.05       |
| 4686 He II  | 1.39±0.03                         | 1.41±0.06                         | 4.36±0.08       |
| 4702 [Fe III] | 0.12±0.01                         | 0.12±0.05                         | 0.39±0.05       |
| 4711 [Ar IV] | 1.08±0.02                         | 1.09±0.06                         | 3.34±0.06       |
| 4713 He I   | 0.70±0.02                         | 0.70±0.05                         | 2.14±0.06       |
| 4734 [Fe III] | 0.04±0.01                         | 0.04±0.05                         | 0.13±0.04       |
| 4740 [Ar IV] | 0.95±0.03                         | 0.96±0.06                         | 3.04±0.08       |
| 4861 H\beta | 100.00±1.43                       | 100.00±1.44                       | 321.70±0.46     |
| 4881 [Fe III] | 0.10±0.02                         | 0.10±0.05                         | 0.31±0.06       |
| 4893 [Fe VII] | 0.04±0.01                         | 0.04±0.05                         | 0.14±0.05       |
| 4907 [Fe IV] | 0.11±0.02                         | 0.11±0.05                         | 0.41±0.05       |
| 4921 He I   | 1.04±0.03                         | 1.03±0.05                         | 4.05±0.11       |
| 4930 [Fe III] | 0.09±0.02                         | 0.09±0.04                         | 0.34±0.06       |
| 4959 [O III] | 107.82±1.54                       | 106.45±1.53                       | 406.70±0.53     |
| 4972 [Fe VI] | 0.11±0.02                         | 0.10±0.04                         | 0.41±0.07       |
| 4986 [Fe III] | 0.41±0.03                         | 0.40±0.05                         | 1.48±0.08       |
| 5007 [O III] | 333.72±4.75                       | 327.89±4.70                       | 1075.00±0.75    |

LR4 ($\lambda$5010–5831)

| Line   | $100 \times F(\lambda)/F(H\beta)$ | $100 \times I(\lambda)/I(H\beta)$ | $EW(\lambda)^a$ |
|--------|-----------------------------------|-----------------------------------|-----------------|
| 4959 [O III] | 107.84±1.54                       | 106.46±1.53                       | 317.50±0.32     |
| 4986 [Fe III] | 0.41±0.02                         | 0.41±0.05                         | 1.37±0.06       |
| 5007 [O III] | 325.05±4.62                       | 319.37±4.57                       | 912.80±0.57     |
| 5016 He I   | 2.01±0.04                         | 1.97±0.06                         | 5.99±0.08       |
| 5041 Si II  | 0.15±0.02                         | 0.14±0.05                         | 0.51±0.05       |
| 5048 He I   | 0.21±0.02                         | 0.21±0.05                         | 0.74±0.07       |
| 5147 [Fe VI] | 0.06±0.03                         | 0.05±0.05                         | 0.19±0.06       |
| 5158 [Fe VII] | 0.05±0.02                         | 0.05±0.05                         | 0.18±0.06       |
| 5176 [Fe VI] | 0.11±0.02                         | 0.11±0.05                         | 0.42±0.08       |
| 5198 [N i]  | 0.21±0.02                         | 0.20±0.05                         | 0.74±0.09       |
| 5200 [N i]  | 0.09±0.02                         | 0.08±0.04                         | 0.31±0.08       |
| 5261 [Fe II] | 0.06±0.02                         | 0.06±0.04                         | 0.21±0.06       |
| 5270 [Fe III] | 0.18±0.03                         | 0.17±0.05                         | 0.64±0.10       |
| 5273 [Fe II] | 0.02±0.01                         | 0.02±0.04                         | 0.09±0.06       |
| 5323 [Cl IV] | 0.05±0.03                         | 0.04±0.05                         | 0.17±0.09       |
| 5335 [Fe VI] | 0.03±0.02                         | 0.03±0.04                         | 0.12±0.09       |
| 5411 He II  | 0.14±0.03                         | 0.13±0.04                         | 0.52±0.11       |
| 5517 [Cl III] | 0.11±0.03                         | 0.10±0.05                         | 0.43±0.11       |
| 5537 [Cl III] | 0.08±0.02                         | 0.07±0.04                         | 0.31±0.11       |
| 5631 [Fe VI] | 0.02±0.02                         | 0.02±0.04                         | 0.08±0.10       |
| 5639 [Fe VI] | 0.03±0.04                         | 0.03±0.05                         | 0.12±0.09       |
| 5721 [Fe VII] | 0.05±0.03                         | 0.04±0.04                         | 0.21±0.13       |

LR5 ($\lambda$5741–6524)

| Line   | $100 \times F(\lambda)/F(H\beta)$ | $100 \times I(\lambda)/I(H\beta)$ | $EW(\lambda)^a$ |
|--------|-----------------------------------|-----------------------------------|-----------------|
| 5755 [N II] | 0.05±0.01                         | 0.05±0.03                         | 0.24±0.03       |
| 5876 He I   | 10.18±0.15                        | 9.25±0.14                         | 47.95±0.13      |
| 5957 Si II  | 0.04±0.01                         | 0.03±0.03                         | 0.19±0.03       |
| 5979 Si II  | 0.06±0.01                         | 0.05±0.03                         | 0.29±0.03       |
| 6046 O I    | 0.05±0.01                         | 0.05±0.03                         | 0.27±0.04       |
| 6102 He II  | 0.04±0.01                         | 0.03±0.03                         | 0.21±0.04       |
| 6118 He II  | 0.01±0.01                         | 0.01±0.02                         | 0.08±0.03       |
| 6133 [Fe III] | 0.02±0.01                         | 0.02±0.02                         | 0.11±0.03       |
| 6300 [O I]  | 0.74±0.02                         | 0.65±0.03                         | 4.10±0.06       |
| 6312 [S III] | 0.64±0.02                         | 0.56±0.03                         | 3.67±0.07       |
| 6348 Si II  | 0.08±0.01                         | 0.07±0.02                         | 0.46±0.04       |
| 6363 [O I]  | 0.27±0.01                         | 0.24±0.02                         | 1.58±0.05       |
| 6371 Si II  | 0.09±0.01                         | 0.08±0.02                         | 0.55±0.05       |
Table 4. Continued.

| Line               | 100×\(F(\lambda)/F(H\beta)\) | 100×\(I(\lambda)/I(H\beta)\) | \(E(W(\lambda)^{a})\) |
|--------------------|---------------------------------|---------------------------------|---------------------------------|
| 6363 [O i]         | 0.28±0.01                       | 0.25±0.03                       | 1.50±0.05                       |
| 6371 Si ii         | 0.09±0.01                       | 0.08±0.03                       | 0.47±0.05                       |
| 6548 [N ii]        | 0.44±0.01                       | 0.38±0.03                       | 1.98±0.05                       |
| 6563 Hα             | 317.90±4.52                     | 274.31±4.26                     | 1427.00±0.77                    |
| 6583 [N ii]        | 1.33±0.02                       | 1.14±0.04                       | 6.23±0.07                       |
| 6678 He i          | 3.07±0.05                       | 2.63±0.05                       | 17.76±0.12                      |
| 6717 [S ii]        | 2.27±0.04                       | 1.94±0.04                       | 13.41±0.11                      |
| 6731 [S ii]        | 1.77±0.03                       | 1.51±0.04                       | 10.49±0.10                      |
| 6739 [Fe iii]      | 0.06±0.01                       | 0.05±0.02                       | 0.33±0.05                       |
| 7002 O i           | 0.07±0.01                       | 0.06±0.02                       | 0.47±0.05                       |
| 7065 He i          | 5.98±0.09                       | 4.99±0.09                       | 38.51±0.19                      |
| 7065 He i          | 5.98±0.09                       | 4.99±0.09                       | 31.85±0.13                      |
| 7136 [Ar iii]      | 1.93±0.03                       | 1.60±0.04                       | 11.48±0.10                      |
| 7171 [Ar iv]       | 0.09±0.01                       | 0.08±0.02                       | 0.59±0.06                       |
| 7237 [Ar iv]       | 0.06±0.01                       | 0.05±0.02                       | 0.37±0.06                       |
| 7254 O i           | 0.11±0.01                       | 0.09±0.02                       | 0.68±0.07                       |
| 7263 [Ar iv]       | 0.04±0.01                       | 0.03±0.02                       | 0.25±0.06                       |
| 7281 He i          | 0.84±0.02                       | 0.70±0.03                       | 5.38±0.10                       |
| 7320 [O ii]        | 0.69±0.02                       | 0.57±0.02                       | 4.47±0.10                       |
| 7330 [O ii]        | 0.61±0.02                       | 0.50±0.02                       | 4.06±0.10                       |
| 7468 N i           | 0.04±0.01                       | 0.03±0.02                       | 0.26±0.07                       |
| 7751 [Ar iii]      | 0.57±0.02                       | 0.46±0.02                       | 4.18±0.12                       |
| 7816 He i          | 0.13±0.01                       | 0.11±0.02                       | 0.96±0.10                       |
| 8046 [Cl iv]       | 0.08±0.01                       | 0.07±0.02                       | 0.64±0.11                       |
| 8223 N i           | 0.06±0.01                       | 0.05±0.02                       | 0.52±0.08                       |
| 8281 P32            | 0.09±0.01                       | 0.14±0.05                       | 0.77±0.09                       |
| 8298 P28            | 0.10±0.01                       | 0.16±0.04                       | 0.90±0.09                       |
| 8306 P27            | 0.09±0.01                       | 0.15±0.05                       | 0.87±0.09                       |
| 8314 P26            | 0.16±0.01                       | 0.20±0.03                       | 1.51±0.10                       |
| 8323 P25            | 0.15±0.01                       | 0.19±0.04                       | 1.38±0.11                       |
| 8334 P24            | 0.28±0.01                       | 0.30±0.03                       | 2.53±0.10                       |
| 8346 P23            | 0.16±0.01                       | 0.20±0.03                       | 1.52±0.09                       |
| 8359 P22            | 0.31±0.01                       | 0.32±0.02                       | 2.89±0.11                       |
| 8362 He i           | 0.14±0.01                       | 0.11±0.01                       | 1.29±0.09                       |
| 8374 P21            | 0.37±0.01                       | 0.35±0.02                       | 3.76±0.13                       |
| 8392 P20            | 0.39±0.01                       | 0.38±0.02                       | 3.53±0.12                       |
| 8413 P19            | 0.38±0.01                       | 0.38±0.02                       | 3.40±0.12                       |
| 8438 P18            | 0.39±0.01                       | 0.38±0.02                       | 3.40±0.11                       |
| 8446 O i            | 0.84±0.02                       | 0.65±0.02                       | 7.11±0.14                       |
| 8467 P17            | 0.58±0.02                       | 0.53±0.02                       | 5.09±0.12                       |
| 8502 P16            | 0.59±0.01                       | 0.54±0.02                       | 5.01±0.14                       |
| 8545 P15            | 0.75±0.02                       | 0.67±0.03                       | 5.62±0.11                       |
| 8598 P14            | 0.88±0.02                       | 0.77±0.03                       | 6.73±0.13                       |
| 8664 P13            | 1.03±0.02                       | 0.89±0.03                       | 7.07±0.11                       |
| 8751 P12            | 1.43±0.03                       | 1.18±0.03                       | 11.98±0.15                      |
| 8863 P11            | 1.72±0.03                       | 1.40±0.04                       | 12.72±0.13                      |
| 9015 P10            | 1.59±0.03                       | 1.29±0.03                       | 13.61±0.15                      |
| 9069 [S iii]        | 4.74±0.07                       | 3.58±0.07                       | 44.31±0.25                      |

\(C(H\beta)\) \(0.190±0.018\)

\(EW_{\text{obs}}(\AA)\) \(1.3±0.2\)

\(F(H\beta)(10^{-14} \text{ erg s}^{-1} \text{ cm}^{-2})\) \(4.3±0.2\)
### Table 5. Emission Line Fluxes and Equivalent Widths in the Integrated Spectrum of the H ii Region

| Line        | $F(\lambda)^a$   | $100 \times F(\lambda)/F(\text{H} \beta)^a$ | $I(\lambda)^b$ | $100 \times I(\lambda)/I(\text{H} \beta)^b$ | $EW(\lambda)^b$ |
|-------------|------------------|---------------------------------------------|----------------|---------------------------------------------|-----------------|
| [O ii] 3727 | 228.1±7.1        | 21.87±0.68                                  | 468.0±14.9     | 25.82±0.82                                  | 23.95±0.45      |
| H10        | 49.1±4.4         | 4.71±0.42                                   | 114.4±12.1     | 6.31±0.67                                   | 6.12±0.40       |
| H9         | 60.4±6.3         | 5.79±0.60                                   | 138.1±16.3     | 7.62±0.90                                   | 6.67±0.62       |
| [Ne iii] 3868 | 218.9±10.2      | 20.99±0.98                                  | 438.4±20.7     | 24.19±1.14                                  | 24.21±1.10      |
| H8 + H2     | 175.8±13.4       | 16.86±1.28                                  | 365.8±28.8     | 20.18±1.59                                  | 21.73±1.20      |
| [Ne iii] 3968 | 225.0±3.8       | 21.57±0.36                                  | 458.4±10.3     | 25.29±0.57                                  | 26.58±0.20      |
| He i 4026   | 15.3±1.3         | 1.47±0.12                                   | 29.9±2.4       | 1.65±0.13                                   | 1.73±0.13       |
| Hδ 4101     | 268.3±4.4        | 25.72±0.42                                  | 532.9±11.4     | 29.40±0.63                                  | 30.79±0.24      |
| Hγ 4340     | 462.8±7.7        | 44.37±0.74                                  | 870.9±15.4     | 48.05±0.85                                  | 79.26±0.57      |
| [O iii] 4363 | 102.4±3.7        | 9.82±0.35                                   | 190.0±6.9      | 10.48±0.38                                  | 19.32±0.51      |
| He i 4471   | 35.8±1.5         | 3.43±0.14                                   | 65.3±2.7       | 3.60±0.15                                   | 6.57±0.24       |
| He ii 4686  | 36.9±1.8         | 3.54±0.17                                   | 65.4±3.3       | 3.61±0.18                                   | 6.70±0.26       |
| [Ar iv] 4740 | 9.3±1.1          | 0.89±0.11                                   | 16.3±2.0       | 0.90±0.11                                   | 1.76±0.24       |
| Hβ 4861     | 1043.0±15.3      | 100.00±1.47                                 | 1812.5±27.0    | 100.00±1.49                                 | 213.40±0.61     |
| He i 4921   | 12.0±1.6         | 1.15±0.15                                   | 20.5±2.7       | 1.13±0.15                                   | 2.58±0.41       |
| [O iii] 4959 | 1078.0±15.6      | 103.36±1.50                                 | 1842.0±26.8    | 101.63±1.48                                 | 204.80±0.27     |
| [O iii] 5007 | 3214.0±46.4      | 308.15±4.45                                 | 5458.5±79.4    | 301.16±4.38                                 | 576.20±0.34     |
| He i 5876   | 115.9±1.8        | 11.11±0.17                                  | 178.2±2.9      | 9.83±0.16                                   | 27.80±0.16      |
| [O i] 6300   | 8.6±0.6          | 0.82±0.06                                   | 12.5±0.9       | 0.69±0.05                                   | 2.36±0.15       |
| [S iii] 6312 | 6.6±0.5          | 0.63±0.05                                   | 9.6±0.9        | 0.53±0.05                                   | 1.69±0.12       |
| Hα 6563     | 3447.9±49.8      | 330.58±4.77                                 | 4972.8±78.5    | 274.36±4.33                                 | 689.40±0.30     |
| [N ii] 6583 | 12.4±0.5         | 1.19±0.05                                   | 17.9±0.7       | 0.99±0.04                                   | 2.48±0.09       |
| He i 6678   | 31.8±0.7         | 3.05±0.07                                   | 45.5±1.1       | 2.51±0.06                                   | 7.46±0.15       |
| [S ii] 6717 | 26.8±0.6         | 2.57±0.07                                   | 38.1±1.0       | 2.10±0.05                                   | 6.38±0.14       |
| [S ii] 6731 | 21.8±0.7         | 2.09±0.07                                   | 30.8±1.1       | 1.70±0.06                                   | 5.22±0.16       |
| He i 7065   | 49.1±1.0         | 4.71±0.10                                   | 68.0±1.6       | 3.75±0.09                                   | 18.59±0.28      |
| [Ar iii] 7136 | 21.1±0.8        | 2.02±0.08                                   | 29.0±1.3       | 1.60±0.07                                   | 7.62±0.29       |
| He i 7281   | 7.6±0.7          | 0.73±0.07                                   | 10.3±0.9       | 0.57±0.05                                   | 2.81±0.29       |

$C(\text{H} \beta)$: 0.240±0.019

$EW_{abs}$: 0.9±0.4

$^a$In units 10^{-16} erg s^{-1} cm^{-2}.

$^b$In Å.