Towards a Complete Census of AGNs in Nearby Galaxies: The Incidence of Growing Black Holes

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Abstract

We investigate the local supermassive black hole (SMBH) density function and relative mass accretion rates of all active galactic nuclei (AGNs) identified in a volume-limited sample of infrared (IR) bright galaxies ($L_{\text{IR}} > 3 \times 10^{9} L_\odot$) to $D < 15$ Mpc (Goulding & Alexander 2009). A database of accurate SMBH mass ($M_{\text{BH}}$) estimates is compiled from literature sources using physically motivated AGN modeling techniques (reverberation mapping, maser mapping and gas kinematics) and well-established SMBH–stellar velocity dispersion relation (hereafter, $\sigma$–$M_{\text{BH}}$–$L_{\text{K, bul}}$ relations). For the three sources without previously published $M_{\text{BH}}$ estimates, we use 2MASS $K$-band imaging and Galfit to constrain the bulge luminosities, and hence SMBH masses. In general, we find the AGNs in the sample host SMBHs which are spread over a wide mass range ($M_{\text{BH}} \approx (0.1–30) \times 10^{6} M_\odot$), but with the majority in the poorly studied $M_{\text{BH}} \approx 10^{8}–10^{9} M_\odot$ region. Using sensitive hard X-ray (2–10 keV) and mid-IR constraints we calculate the bolometric luminosities of the AGNs ($L_{\text{Bol, AGN}}$) and use them to estimate relative mass accretion rates. We use these data to calculate the volume-average SMBH growth rate of galaxies in the local Universe and find that the AGNs hosting SMBHs in the mass range $M_{\text{BH}} \approx 10^{6}–10^{7} M_\odot$ are dominated by optically unidentified AGNs. These relatively small SMBHs are acquiring a significant proportion of their mass in the present-day, and are amongst the most rapidly growing in the local Universe (SMBH mass doubling times of $\approx 6$ Gyr). Additionally, we find tentative evidence for an increasing volume-weighted AGN fraction with decreasing SMBH mass in the $M_{\text{BH}} \approx 10^{8}–10^{9} M_\odot$ range. Overall, we conclude that significant mass accretion onto small SMBHs may be missed in even the most sensitive optical surveys due to absent or weak optical AGN signatures.

Key words: galaxies: active – galaxies: evolution – galaxies: nuclei – infrared: galaxies

1 INTRODUCTION

It is now well established that all massive galaxies ($M_* \approx 10^{10}–10^{12} M_\odot$) in the local Universe harbour central supermassive black holes (SMBHs), with masses proportional to those of their stellar spheroids (hereafter, bulge; e.g., Kormendy & Richstone 1995; Magorrian et al. 1998). Comparisons between the SMBH mass density in the local Universe and the total energy produced by active galactic nuclei (AGNs) across cosmic time have shown that these SMBHs were primarily grown through mass-accretion events (e.g., Soltan 1982; Rees 1984; Marconi et al. 2004). The space density of high-luminosity AGNs appears to have peaked at higher redshifts than lower-luminosity AGNs, suggesting that the most massive SMBHs ($M_{\text{BH}} \approx 10^{8}–10^{9} M_\odot$) grew first, a result commonly referred to as ‘AGN cosmic downsizing’ (e.g., Cowie et al. 2003; Ueda et al. 2003; McLure & Dunlop 2004; Hasinger et al. 2004; Alonso-Herrero et al. 2008). Extrapolation of these results imply that the most rapidly growing SMBHs in the nearby Universe should be of comparatively low mass ($M_{\text{BH}} \ll 10^{8} M_\odot$). To determine the characteristic masses of these growing SMBHs requires a complete census of AGN activity and SMBH masses in the local Universe.

Using data from the Sloan Digital Sky Survey (SDSS; York et al. 2000) in conjunction with the well established SMBH–stellar velocity dispersion relation (hereafter, $\sigma$–$M_{\text{BH}}$; e.g., Gebhardt et al. 2000; Tremaine et al. 2002; Heckman et al. 2004; hereafter, H04) deduced that relatively low mass SMBHs ($M_{\text{BH}} \approx 3 \times 10^{7} M_\odot$) residing in

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moderately massive bulge-dominated galaxies host the majority of present-day accretion onto SMBHs. However, the space density of SMBHs derived from the M—σ relation in the optical survey of H04 was limited by the spectral resolution of the SDSS (σ* > 70 km s^{-1}) to SMBHs of M_{BH} ≳ 3 × 10^6 M_⊙ (assuming the M—σ relation of Gebhardt et al. 2000). Furthermore, due to attenuation of optical emission by dust, source selection and AGN classification at optical wavelengths will be biased against gas-rich, dust-obscured objects. These surveys are unlikely to include galaxies hosting the smallest bulges, and consequently the lowest mass SMBHs, and may therefore be missing a significant proportion of SMBH growth in the local Universe. Indeed, the nearby Scd galaxy, NGC 4945, hosting a low-mass SMBH (M_{BH} ≈ 1.4 × 10^6 M_⊙; [Greenhill et al. 1997]) only displays evidence for AGN activity in X-ray (Iwasawa et al. 1993) and mid-infrared (mid-IR) observations (Goulding & Alexander 2009). By contrast, the AGN in NGC 4945 (accreting at ≳ 30 percent of the predicted Eddington limit; [Itoh et al. 2008]) is completely hidden at optical wavelengths, and classified as a starburst galaxy. Clearly, using optical data alone, the intrinsic AGN properties of sources similar to NGC 4945 cannot be derived.

While optical emission-line diagnostics alone cannot reliably characterise the properties of a non-negligible fraction of the AGN population, they are readily identified at obscuration independent wavelengths (e.g., X-ray; mid-IR). Hence, the identification of AGNs made at X-ray and mid-IR wavelengths complements traditional UV/optical methods to yield a more complete census of AGN activity. Indeed, using the high resolution mid-IR spectrograph onboard the NASA Spitzer Space Telescope (Spitzer-IRS; [Goulding & Alexander 2009; GA09]) found using the first complete volume-limited sample of all (≈ 94 percent) local (D < 15 Mpc) bolometrically luminous galaxies (L_{IR} > 3 × 10^9 L_⊙), that ≈ 50 percent of local AGNs are not identified in sensitive optical surveys. At least 30 percent of these AGNs were previously identified as pure optical starburst galaxies, similar to NGC 4945 (i.e., not even otherwise known to be transition-type objects as defined by [Kauffmann et al. 2003]). Furthermore, ≈ 30 percent of the optically unidentified AGNs were found to reside in late-type spiral galaxies (Sc–Sd; e.g., similar to NGC 4945). Complimentary to this, from a heterogenous sample of Palomar galaxies, [Satyapal et al. 2007] and [Satyapal et al. 2008] have also concluded that optically unidentified AGNs exist in some late-type spiral galaxies. With the inclusion of these new optically unidentified AGNs, it is natural to ask, what are the masses of local active SMBHs, what are their Eddington ratios, and hence, how rapidly are active SMBHs growing in the local Universe?

In this paper, we investigate the growth rates and space density of actively accreting SMBHs using the 17 AGNs identified in the volume-limited survey of GA09. Whilst the source statistics considered here are significantly smaller than those studies using the SDSS, this work compliments that of H04 by including a relatively large number (given the considered small volume) of optically unidentified AGNs (10) which would not be reliably identified or characterised in the SDSS survey. Furthermore, by including a significant population of bolometrically luminous (but dust-obscured) late-type spiral galaxies (Sc–Sd) we are able to extend the SMBH density function to M_{BH} < 3 × 10^6 M_⊙. As many of the late-type spiral galaxies host small galactic bulges, and hence lower mass SMBHs, particular attention is paid to obtaining accurate mass estimates for these SMBHs. Given their proximity, many of the sources in GA09 are well-studied and have multiple estimates of SMBH mass (M_{BH}) from a variety of methods (i.e., reverberation mapping techniques; mapping of water maser spots; gas kinematical estimates; the M—σ relation; correlation of M_{BH} with the luminosity of the galactic bulge); below we discuss the relative accuracy of each SMBH mass estimate technique. Furthermore, to determine the relative mass accretion rates and hence average growth times of the SMBHs in our sample we require the best available estimates of the AGN bolometric luminosity (L_{Bol,AGN}). Here we use two approaches: 1) for the AGNs with currently published data, we use high-quality well-constrained sensitive hard X-ray (2–10 keV) luminosities to directly measure L_{Bol,AGN}; and 2) we accurately infer L_{Bol,AGN} using a well-constrained hard X-ray to high-ionisation mid-IR emission line relation.

In §2 we outline the construction and basic reduction analysis of the AGN sample derived from GA09. In §3 we present the SMBH mass estimates. For a minority of objects (three out of 17 AGNs) without published M_{BH} estimates we outline the use of a bulge/disc decomposition method with 2MASS K-band images, and following Marconi & Hunt (2003), we use the M_{BH}—L_{bol} relation to estimate their SMBH masses. In §4 we use hard (2–10 keV) X-ray measurements and high-ionisation mid-IR emission to estimate the intrinsic luminosity of the AGNs considered in our sample. Using our well-defined estimates for SMBH mass and AGN bolometric luminosity, we investigate the relative mass accretion rates of our sample of active SMBHs in §5. We use these estimates to provide new constraints on the volume-average SMBH growth rates in the local Universe. We further compare these results to the previous works of H04 and [Greene & Ha (2007)] by producing a local AGN population density function. Finally, in §6 we present our conclusions.

2 THE SAMPLE

2.1 Selection and Data-reduction

The sample of local AGNs is derived from the Spitzer-IRS spectral investigation of a volume-limited sample of IR-bright galaxies to D < 15 Mpc by GA09. High-resolution Spitzer-IRS spectroscopy (R ~ 600) was obtained for sixty-four of the sixty-eight (≈ 94 percent) galaxies detected in the Revised Bright Galaxy Sample (RBGS; [Sanders et al.])
with \( L_{\text{IR}} \gtrsim 3 \times 10^9 L_\odot \) within the considered volume. A particular advantage to a volume-limited sample is that, unlike magnitude-limited surveys, they do not suffer from radial selection effects and thus can be used to construct volume-averaged statistics. By contrast, volume-limited samples yield limited source number statistics due to their inability to probe both the faintest and most luminous systems.

Given the relatively small distance scale considered here, it is prudent to note the importance of distance measurement. For our consider sample, luminosity distances were calculated using the cosmic attractor model of Mould et al. (2000) which accurately adjusts heliocentric redshifts to the centroid of the local group. Hence, unlike many other local surveys, to \( D < 15 \) Mpc our sample does not include galaxies from local over densities such as the Virgo cluster at \( D \sim 16 \) Mpc (see Appendix A1 for a detailed analysis and validation of the considered space-volume in this survey).

The Spitzer-IRS spectroscopic data presented in GA09 was reduced using a custom IDL pipeline which utilises the Spitzer Science Center data-processing packages SPICE, IRSCLEAN and CUBISM. For further detailed information on the reduction processes and spectral analyses see §§2.1 and 2.2 of GA09 and references there-in.

### 2.2 [NeV] as an unambiguous AGN indicator

Due to the very high-ionization potential of [NeV] (97.1 eV) we consider its detection coincident with the galactic nucleus in mid-IR spectroscopy to be an almost unambiguous identifier of AGN activity. Theoretically, Schauer & Stasińska (1999) have predicted that extremely hot O and B stars, in particular dense populations of Wolf-Rayet stars, may produce ionization spectra capable of exciting lines such as [OIV] (54.9 eV) and [NeV]. However, observationally, GA09 found that even in extreme Wolf-Rayet galaxies, whilst [OIV] is clearly detected in these types of systems, [NeV] emission remains absent to the detection limits of this survey. Complimentary to this, Hao et al. (2009) find from a Spitzer-IRS study of 12 Blue Compact Dwarf galaxies that the mid-IR spectroscopy for eight of their sample contain [OIV] emission, however none appear to be producing [NeV].

Similarly, extreme starburst driven shocks have also been predicted to excite some high-ionization lines such as [NeV] (Allen et al. 2008), however these require exceptionally high velocities, and based on the [NeV]\( \lambda \)14.32 \( \mu \)m-[NeII]\( \lambda \)12.81 \( \mu \)m and [NeIII]\( \lambda \)15.51 \( \mu \)m-[NeII]\( \lambda \)12.81 \( \mu \)m emission line ratios presented in GA09, the AGNs in this study are not consistent with shock models. Furthermore, AGNs which contain strong star-formation contributions to their bolometric luminosity are often found to have relatively low mid-IR [NeV]-[NeII] and [NeIII]-[NeII] ratios (Armus et al. 2006; Satyapal et al. 2008; Goulding & Alexander 2009; Dale et al. 2009). However, these ratios are not necessarily strong tracers of the intrinsic power of the AGN. Indeed, many Seyfert 2s are found to be hosted in galaxies where-by star-formation dominates the IR spectral energy distribution (Weedman et al. 2003; Buchanan et al. 2006; Deco et al. 2007), thus yielding a low [NeV]-[NeII] ratio (log-average \( \approx 0.02 \)); however, the central source may still be extremely luminous at other energies e.g., NGC 4945 is the most luminous local AGN at \( E > 20 \) keV (Done et al. 1996) and by contrast has a [NeV]-[NeII] ratio of \( \approx 0.01 \).

The mid-IR spectra of seventeen of the sixty-four galaxies (\( \approx 27 \pm 6 \) percent) presented in GA09 were found to contain the [NeV] \( \lambda \)14.32 \( \mu \)m emission line, and hence, host AGN activity. These seventeen sources are the main focus of the current paper. See Table 1.

### 3 BLACK HOLE MASS DETERMINATION

#### 3.1 Archival Data

To accurately determine the relative mass accretion rates and space density of active SMBHs in the local Universe requires reliable SMBH mass (\( M_{\text{BH}} \)) estimates for the seventeen AGNs within our volume-limited sample. Here we outline the construction of the heterogeneous database of the most reliable available SMBH masses for these AGNs, derived from a variety of archival sources and \( M_{\text{BH}} \) estimation methods.

Many of the sources in the sample are late-type galaxies hosting relatively small bulges, and hence low-mass SMBHs. In such systems, SMBH mass estimates are often challenging to determine as: 1) characteristically low velocity dispersions can be difficult to measure as they are often at the resolution limit of published observations, 2) modeling of the contamination from composite stellar populations can often lead to inconsistencies between published measurements and 3) SMBH mass relations are poorly constrained at \( M_{\text{BH}} \sim 10^6 \) \( M_\odot \).

Given the varying degrees of accuracy associated with \( M_{\text{BH}} \) measurements estimated from differing methodologies, we have chosen to prioritise the archival data (which we further expand on here) based upon two broad categories: (i) physically motivated AGN modeling techniques (i.e., reverberation mapping, water maser mapping and gas kinematics of the central engine) and (ii) indirect estimations from observational relations (the \( M-\sigma \) and \( M_{\text{BH}}-L_{\text{bul}} \) relations). The adopted \( M_{\text{BH}} \) for the sources and their associated measurement methodologies are given in columns 12 and 13 of Table 1. The database contains four SMBH measurements determined from physically motivated AGN modeling and 13 from indirect methods.

#### 3.1.1 Direct SMBH mass constraints from reverberation mapping, maser mapping and gas kinematics

Under the assumption that the gas in the broad-line region (BLR) is virialised by the SMBH and the orbital motion of the gas is Keplerian, \( M_{\text{BH}} \) estimations are possible through reverberation mapping techniques (Blandford & McKee 1982). The time-lag between changes
in the AGN continuum flux and the response of the broad-line region is used to directly infer the size of the virial radius and hence the mass of the SMBH (for an in-depth review see Peterson 2001). To date, reverberation mapping is widely accepted to be the most reliable of M_{BH} estimation methodologies (e.g., Wandel et al. 1999; Peterson & Wandel 1999; Onken et al. 2003; Bentz et al. 2009). Only a minority of the galaxies in the sample (NGC 1068, NGC 4051 and NGC 5033) are known to have detected BLRs. However, the BLRs of NGC 1068 and NGC 5033 (both are Seyfert 1.9) are extremely weak and are therefore likely to suffer from optical extinction. Thus, M_{BH} estimates using the BLR are not possible for these two AGNs. Hence, only one object (NGC 4051; Wandel 1999) in our sample has archival reverberation mapping data giving M_{BH} \approx 1.4^{+1.5}_{-0.9} \times 10^6 M_\odot.

Complimentary to reverberation mapping, and also assuming Keplerian motion, mapping of water maser spots and observations of gas kinematics within the gravitational sphere of influence of the SMBH are thought to yield relatively precise measurements of M_{BH}. Again, due to the nature of the observations and the requirement of a suitable gas disc (i.e., a relatively face-on inclination to the observer), few M_{BH} estimations using these methods exist in the current literature. Indeed, in our sample there are currently only two AGNs (NGC 1068 and NGC 4945) with M_{BH} estimations from the mapping of water maser spots (M_{BH} \approx 1.6 \times 10^6 and 1.1 \times 10^6 M_\odot, respectively) and one AGN (NGC 5128) with a spatially resolved gas dynamical M_{BH} estimate (M_{BH} \approx 2.4 \times 10^8 M_\odot).

3.1.2 Indirect SMBH mass constraints from the M–σ relation

Since the seminal discovery that the mass of the stellar spheroid is closely related to M_{BH} (Magorrian et al. 1998), indirect M_{BH} estimation methods have become ubiquitous in the current literature (e.g., Ferrarese & Merritt 2000; Gebhardt et al. 2000; Gebhardt et al. 2003; Onken et al. 2004; Greene & Ho 2006). Comparisons of spatially resolved stellar dynamics with reverberation-based M_{BH} estimates (over the range of M_{BH} \approx 2 \times 10^5–10^8 M_\odot) show that these are well-correlated with the effective stellar velocity dispersions (σ_e) within the galaxy bulge (i.e., the M–σ_e relation; Gebhardt et al. 2000).

\[ M_{BH} = (1.2 \pm 0.2) \times 10^8 M_\odot \left( \frac{\sigma_e}{200 \text{ km s}^{-1}} \right)^{3.75 \pm 0.30} \]  

with an intrinsic scatter of only 0.34 dex. Furthermore, with the inclusion of a sample of dwarf Seyfert 1 galaxies, the correlation may be reasonably extrapolated to SMBH masses in the range 10^5–10^8 M_\odot (Barth et al. 2003). However, we acknowledge that to-date, the M–σ_e relation is still poorly constrained for M_{BH} < 10^6 M_\odot and suffers from systematic uncertainties. Specifically, these are caused by the lack of homogeneous source statistics in this mass region and also the poor understanding of the physical nature of the central region in very late-type (Scd–Sm) galaxies (e.g., the existence of either a pseudo- or classical bulge; Kormendy & Kennicutt 2004). Of the three AGNs in our sample which we classified as Hubble-type Scd or later, only NGC 3621 (Sd) has an M_{BH} measurement estimated from the M–σ_e relation. Specifically, using the Echellette spectrograph on Keck-II, Barth et al. (2009) measure an accurate line-of-sight stellar velocity dispersion of the Calcium triplet lines (CaT; λλ 8498, 8542, 8662) observed in the central nuclear star cluster. They find a velocity dispersion consistent with a SMBH mass of M_{BH} \approx 3 \times 10^5 M_\odot. We find that all of the AGNs in our sample host SMBHs with M_{BH} \gtrsim 10^5 M_\odot and thus are unlikely to suffer significant systematic uncertainties arising from the use of the M–σ_e relation to estimate M_{BH} even in the most late-type galaxies.

Ten of the 13 AGNs in our sample without direct M_{BH} measurements have published M_{BH} estimates using

\[ M_{BH} \approx \left( 1.4^{+1.5}_{-0.9} \right) \times 10^6 M_\odot \]

\[ \sigma_e = 200 \text{ km s}^{-1} \]

\[ \left( \frac{\sigma_e}{200 \text{ km s}^{-1}} \right)^{3.75 \pm 0.30} \]
from the Two Micron All-Sky Survey (2MASS) extended source catalogue and consist of pre-mosaicked (1 arcsecond per pixel resolution) all-sky atlas images. The K-band images of the three galaxies were modeled with a central point spread function (PSF) and a constant sky background contribution, whilst the bulge and host-galaxy components were modeled using variations of the Sersic profile:

$$\Sigma(r) = \Sigma_e e^{-\kappa((r/r_e)^{1/n} - 1)}$$

(2)

where $r_e$ is the effective radius of the profile, $\Sigma_e$ is the surface brightness at the effective radius, $n$ is the power-law (Sersic) index, and $\kappa$ is coupled to $n$ such that half of the total flux of the object is within the effective radius. We employ two special forms of the Sersic profile in our galfit modeling, the exponential ($n = 1$) and the de Vaucouleurs ($n = 4$) profiles, which are classically used to model galactic discs and bulges, respectively.

Häussler et al. (2007) have shown that the reliability of the fitting parameters produced by galfit are strongly dependent on the initial estimates. Galfit will, in general, fail to find the overall global chi-squared ($\chi^2$) minimum to the fit if the initial estimates are poorly constrained. Thus, to reduce this systematic effect, and aid the fitting routine, we use a simplified 1-dimensional fit to produce initial estimates of the fitting parameters. A 1-dimensional surface-brightness slice of the K-band image was taken across the major axis of each of the galaxies. A surface brightness profile extending from the nucleus was produced by averaging the two semi-major axes from the slice, and removing the measured background flux. Few spiral galaxies are found to host bulges with true de Vaucouleurs profiles, thus a global $\chi^2$ reduction process was used to simultaneously fit a generalised Sersic profile and a fixed ($n = 1$) exponential disc to the 1-d surface brightness profile. From these, we calculate Sersic and disc radii, as well as the Sersic index of the bulge. Combining the 1-d parameter estimates with the total K-band magnitude from the 2MASS Large Galaxy Atlas (Jarrett et al. 2004), we generate an appropriate set of constraints and initial parameters to be input to galfit.

Using the derived parameter estimates, galfit is used to fit a generalised 2-d Sersic profile with an exponential disc to the K-band image. To again aid the Galfit reduction analysis, particular attention is paid to simulating accurate PSFs for the 2MASS images using known standard stars (J. R. Lucey private communication). Results of this bulge/disc reduction for the three objects are presented in Fig. 1 and column 11 of Table 1.

We have directly tested our robust galfit method using the late-type galaxies (i.e., S0 or later) presented in the dataset of MH03 and find close agreement ($\approx 0.1$ dex). We do note however, that we find a systematic offset of a factor $\approx 2$ in bulge luminosity for the AGNs in MH03 that are hosted in low-inclination angle late-type galaxies, which is likely to be caused by Galfit over-estimating the contribution of the bulge to the total flux of the galaxy. Indeed, when directly comparing a sample of reverberation mapped X-ray detected AGNs to $M_{BH}$ estimations using the MH03 formalism, Vasudevan et al. (2009) find similar results. However,
The three galaxies fitted in our sample are all moderately to highly inclined and thus this systematic effect will be negligible of fitting; however, the bulge fit does not appear to be capable of fitting; yet this systematic effect will be negligible. We note that the well established M–σ relation for AGNs may be systematically over-estimated by a factor of 2 (see above). Thus, our derived Eddington rate for NGC 1792, presented in §5.1, should be considered a lower limit.

From the obtained K-band bulge magnitudes (column 11, Table 1) we calculate bulge luminosities (L_{K,bul}). Using the well established M_{BH} to bulge K-band luminosity relation of MH03 (hereafter, M_{BH}–L_{K,bul}),

\[
\log M_{\text{BH}} = (8.08 \pm 0.10) + (1.21 \pm 0.13)\log \left( \frac{L_{\text{K,bul}}}{8 \times 10^{10}} \right)
\]  

we calculate M_{BH} for ESO 121-G006, NGC 1448 and NGC 1792. The 1-σ uncertainties associated with the M_{BH} values are assessed by combining in quadrature the bulge magnitude errors as calculated by GALFIT with the intrinsic dispersion observed in the M_{BH}–L_{K,bul} relation. The resultant M_{BH} estimates and 1-σ uncertainties are provided in column 12 of Table 1.
4 BOLOMETRIC CORRECTIONS

To determine the relative mass accretion rates, and hence the average growth times of SMBHs in the local Universe we require relatively accurate estimates of the bolometric luminosities of the AGNs ($L_{\text{bol, AGN}}$) in our sample. We use two methods to estimate $L_{\text{bol, AGN}}$: 1) a direct approach using the best available measured hard X-ray (2–10 keV) luminosities, and 2) a well-constrained $L_{\text{bol, AGN}}$–[OIV] luminosity relation to infer the intrinsic luminosity of the AGN (e.g., Dasyra et al. 2008; Meléndez et al. 2008, hereafter, M08).

4.1 Hard X-ray Luminosity as a Tracer of the Bolometric Luminosity of an AGN

High quality hard X-ray spectral analyses arguably provide the most unambiguous method for measuring the intrinsic luminosity of an AGN since: 1) X-rays are relatively unaffected by dust extinction; 2) intrinsic absorption can be directly constrained from high S/N data; 3) star-formation contamination is often found to be negligible. We therefore divide our sample into two categories based on the quality and energy range of their available published X-ray data: 1) AGNs with high S/N X-ray spectra where $N_\text{H}$ has been accurately constrained and/or AGNs with $E > 10$ keV constraints where the observed X-ray emission will only be strongly absorbed for heavily Compton-thick ($N_\text{H} > 10^{25}$ cm$^{-2}$) sources (8 AGNs); and 2) those AGNs with no or low S/N X-ray data, i.e., where there are insufficient counts to accurately determine $N_\text{H}$ and the X-ray flux could have large contributions from star-formation (9 AGNs). Hence, we specifically do not estimate $L_{\text{bol, AGN}}$ for those AGNs with $L_X$ measurements using Chandra that also have no further hard X-ray spectral constraints ($E > 10$ keV) due to the limited band-pass of the instrument at $z \sim 0$ (0.5–8 keV).

For the AGNs in our sample currently with either high S/N spectroscopy or $E > 10$ keV constraints, we estimate $L_{\text{bol, AGN}}$ using the Eddington ratio independent AGN bolometric corrections outlined in Equation 21 of Marconi et al. (2004). We note that Vasudevan & Fabian (2009) have suggested that the bolometric correction factor ($\kappa_{2–10keV} = L_{\text{bol}}/L_{2–10keV}$) may be a function of the Eddington ratio of a considered source. Values of $\kappa_{2–10keV} \sim 10–30$ are considered to be relatively low bolometric corrections and are generally found in AGNs with $\eta < 0.1$ (e.g., Vasudevan & Fabian 2009; Vasudevan et al. 2010). For the sample of X-ray detected AGNs considered here, we calculate similarly consistent values of $\kappa_{2–10keV} \sim 8–30$, and thus conclude that the Eddington ratio is unlikely to be dominating the bolometric corrections adopted here from Marconi et al. (2004).

For further consistency, all archival X-ray luminosities were adjusted to the distances adopted in column 2 of Table 1. Final adopted $L_X$ measurements, estimated $L_{\text{bol, AGN}}$ from $L_X$, and archival references for $L_X$ are given in columns 6, 7 and 9 of Table 1, respectively.

4.2 [OIV] Luminosity as a Tracer of the Bolometric Luminosity of an AGN

For those galaxies in the sample without good-quality hard X-ray constraints we require an additional approach to estimate the intrinsic luminosity of the AGN, and hence the relative mass-accretion rate. Here we build-upon an $L_{\text{bol, AGN}}$ estimation which relies on the AGN-produced [OIV] $\lambda 25.89$ µm luminosities ($L_{\text{OIV}}$) for the AGNs in our sample (e.g., Dasyra et al. 2008; M08).

Based on the simplest Unified Model of AGN (Antonucci 1993), the hot dust within the predicted torus, close to the central engine, reprocesses absorbed UV, optical and X-ray emission into mid-IR emission. Hence, AGN emission detected at IR wavelengths is likely to be isotropic and independent of viewing angle. As discussed in §2, the detection of high-ionisation [NeV] emission (97.1 eV) coincident with the nucleus of a galaxy is considered a robust indicator of AGN activity (e.g., Armus et al. 2006; GA09). Complimentary to this, GA09 find that [NeV] emission is also well correlated with [OIV] emission (54.9 eV) with an intrinsic scatter of only 0.24 dex. As [NeV] emission, and thus [OIV] emission, do not suffer from significant star-formation contamination and are both comparatively extinction-free, they may be used as relatively clean proxies for the bolometric luminosity of the AGNs presented here ($L_{\text{bol, AGN}}$).
Table 2. Catalogue of revised 2–10 keV luminosities for a subset of M08 sample.

| Common Name | log($L_{2-10keV,M08}$) | log($L_{2-10keV,alt.}$) | C-thick? | Ref. | AGN? |
|-------------|--------------------------|--------------------------|----------|------|------|
| Circinus    | 40.58                    | 42.04                    | √        | 1    |      |
| Mrk 3       | 41.95                    | 43.20                    | √        | 2    |      |
| NGC 1365    | 40.99                    | 42.40                    | √        | 3    |      |
| NGC 2992    | 41.69                    | 42.50                    | ×        | 4    |      |
| NGC 3079    | 40.02                    | 42.18                    | √        | 5    |      |
| NGC 4388    | 41.91                    | 42.57                    | ×        | 6    |      |
| NGC 6240    | 42.23                    | 44.00                    | √        | 7    |      |

Notes: (1) Common galaxy name. (2) Logarithm of 2–10 keV luminosity adopted by M08. (3) Logarithm of absorption corrected 2–10 keV luminosity adopted in this work from individual studies. (4) Is the AGN a Compton thick source? (5) Reference for adopted $L_{2-10keV}$ measurement.

References: (1) Yang et al. (2009); (2) Awaki et al. (2008); (3) Risaliti et al. (2009); (4) Yaqoob et al. (2007); (5) Iyomoto et al. (2001); (6) Shirai et al. (2008); (7) Vignati et al. (1990).

Indeed, for a sample of 35 well-studied optically unobscured AGNs, Dasyra et al. (2008) show that both [NeV] and [OIV] emission are well correlated with the luminosity of the 5100 Å optical continuum, and hence $L_{bol,AGN}$ with an intrinsic scatter of 0.46 and 0.47 dex, respectively. However, the relation of Dasyra et al. (2008) is derived from AGNs with $L_{[OIV]} > 2 \times 10^{40}$ erg s$^{-1}$ ($L_{[NeV]} > 7 \times 10^{39}$ erg s$^{-1}$), and hence we test whether it may be reliably extrapolated to the more modest luminosity AGNs considered here (log-average $L_{[OIV]} \approx 2 \times 10^{38}$ erg s$^{-1}$).

We combine our robustly adopted $L_{bol,AGN}$ from 2-10 keV flux measurements with the X-ray catalogue of nearby ($z < 0.08$) Seyfert galaxies in the Swift-BAT survey which have published [OIV] luminosities in M08. The catalogue of sources in M08 contain 2-10 keV luminosities obtained primarily from ASCA data, 14-195 keV luminosities from the Swift-BAT survey and [OIV] luminosities from Spitzer-IRS spectroscopy. Combining the M08 sample with our 8 AGNs with high-quality X-ray constraints, the range covered in $L_{[OIV]}$ is $\approx (0.7-7000) \times 10^{39}$ erg s$^{-1}$. For consistency, we convert the M08 2-10 keV luminosities to $L_{bol,AGN}$ using the same bolometric corrections adopted in §4.1. In Fig. 2 we plot $L_{bol,AGN}$ versus $L_{[OIV]}$ for the M08 sample and the 8 AGNs in our GA09 sample with good-quality hard X-ray constraints (grey filled squares and blue filled squares, respectively).

Seven of the AGNs in M08 were found to have unusually low 2–10 keV luminosities when compared to the quoted 14–195 keV luminosities, which are highlighted with open circles in Fig. 2 and are shown in Table 2. Using high-quality X-ray spectral analyses, in the literature we find that five of these AGNs are identified as Compton thick AGNs (see column 4 of Table 2). This additional obscuration does not appear to have been accounted for in the adopted value of M08, and therefore we have chosen to select measurements of $L_{2-10keV}$ from the literature, which now give good agreement with 14–195 keV luminosities (see columns 2, 3 and 5 of Table 2; dotted-lines in Fig. 2). The other two AGNs (NGC 2992 and 4388) in M08 with inconsistent 2–10 and 14–195 keV fluxes are found to be highly variable (Beckmann et al. 2007 and Elvis et al. 2004, respectively).

With the inclusion of the obscuration-corrected 2–10 keV luminosities to infer $L_{bol,AGN}$, we find a strong correlation between $L_{[OIV]}$ and $L_{bol,AGN}$ which is characterised by the equation:

$$\log \left( \frac{L_{bol,AGN}}{10^{44} \text{erg s}^{-1}} \right) = (0.38 \pm 0.09) \log \left( \frac{L_{[OIV]}}{10^{44} \text{erg s}^{-1}} \right) + (1.31 \pm 0.09)$$

with an intrinsic scatter in the data of $\approx 0.35$ dex. We find good agreement with the Dasyra et al. (2008) relation (dashed-line) in the region $L_{[OIV]} \approx 10^{40} - 10^{43}$ erg s$^{-1}$ (i.e., where the Dasyra et al. (2008) relation is well sampled). However, for $L_{[OIV]} < 10^{40}$ erg s$^{-1}$ we show that the Dasyra et al. (2008) relation will over estimate $L_{bol,AGN}$ by typically $\approx 0.5$ dex.

In Fig. 2 we highlight two AGNs from the $D < 15$ Mpc sample (NGC 4945 and 5643) which appear to be significant outliers of the observed correlation. NGC 5643 possibly harbours a variable central source. From the detection of strong Fe Kα emission, Maiolino et al. (1998) suggest from using BeppoSAX data that NGC 5643 is possibly Compton thick ($N_H \approx 10^{25}$ cm$^{-2}$); however, Guainazzi et al. (2004) find using XMM-Newton data, that it may be Compton thin with $N_H \approx (6-10) \times 10^{23}$ cm$^{-2}$. Thus, from current available data, the true intrinsic luminosity of the AGN is highly uncertain. Here we conservatively adopt the Compton thin $L_X$ value of Guainazzi et al. (2004); however, we note that if we use the value of Maiolino et al. (1998), then NGC 5643 would lie on our derived relationship. NGC 4945, by contrast, has observations using Ginga, ASCA, OSSE, the Rossi X-ray Timing Explorer, BeppoSAX, Suzaku and most recently Swift-BAT (Iwasawa et al. 1993, Tanaka et al. 1994, Done et al. 1996, Madejski et al. 2000, Guainazzi et al. 2000, Itoh et al. 2008 and Tueller et al. 2008, respectively) all of which provide excellent and consistent spectral constraints over a wide X-ray band (1–200 keV) showing that the central source is Compton thick ($N_H \approx 4 \times 10^{24}$ cm$^{-2}$), with $L_{X,2-10keV} \approx (3-10) \times 10^{42}$ erg s$^{-1}$ and an observed intrinsic variability of a factor $\approx 2$. Here we adopt the luminosity from the most recent observation by Swift, $L_X \approx 3 \times 10^{42}$ erg s$^{-1}$. On the basis of our mid-IR constraints, it would therefore appear that NGC 4945 is under-luminous in [OIV] flux by a factor of $\approx 50$ (see Table 1 and Fig. 2). We suggest this deficit in observed [OIV] flux is unlikely to be due to host galaxy extinction; the required absorption to account for a factor of $\approx 50$ flux difference is $A_V \approx 240$ mags ($N_H \approx 5 \times 10^{23}$ cm$^{-2}$ assuming typical dust-to-gas ratios; using $A_V/E(B-V) = 3.1$). Another possible explanation is a temporary decoupling of the X-ray emitting and narrow-line regions (i.e., the highly-luminous state of NGC 4945 may be a somewhat recent event). Given the spatial difference
Figure 3. AGN bolometric luminosity ($L_{\text{bol,AGN}}$ in ergs s$^{-1}$) is plotted against SMBH mass ($M_{\text{BH}}$) for the $D < 15$ Mpc mid-infrared identified AGNs presented in GA09. Associated 1–σ error bars for $M_{\text{BH}}$ and $L_{\text{bol,AGN}}$ estimations are shown (see §3 and §5.1, respectively for details of their derivations). AGNs which are previously identified in optical surveys are highlighted with open circles. AGNs with $M_{\text{BH}}$ estimates from reverberation mapping (downward triangles), gas dynamics (hour glass), maser mapping (diamond), the $M$–$\sigma_*$ relation (squares) and the $M_{\text{BH}}$–$L_{K,\text{bol}}$ relation (upward triangles) are plotted. NGC 5195 is represented with both an upward triangle and square (with a dashed-line connector) as both of the $M_{\text{BH}}$ estimates for this galaxy are highly uncertain given its irregular morphology.

(And hence, the light-travel time) between the two emission regions: $10^{-4}$ pc and 1–10 pc, respectively, the photoionisation of the narrow-line region, and thus the observed $[\text{OIV}]$ emission, may take $\approx 100$ years to respond to the changes in the X-ray emitting region. We therefore suggest that the intrinsic scatter in the observed $[\text{OIV}]$–$L_{\text{bol,AGN}}$ relation may be significantly reduced if it was possible to account for variability in the central region of all of the AGNs. Indeed, we find the average dispersion decreases to $\approx 0.31$ dex if we remove NGC 4945 from our analysis.

Using Equation 4 and $L_{[\text{OIV}]}$ from GA09 (column 5 of Table 1), we estimate $L_{\text{bol,AGN}}$ (column 7 of Table 1) for those AGNs in our sample currently without good hard X-ray measurements. We use these $L_{\text{bol,AGN}}$ estimates to assess the relative mass accretion rates ($\eta \sim L_{\text{bol,AGN}}/L_{\text{Edd}}$) of the SMBHs in our $D < 15$ Mpc sample.
5 RESULTS AND DISCUSSION

From a volume-limited sample of 64 bolometrically luminous galaxies to $D < 15$ Mpc ($\approx 94$ percent complete) GA09 unambiguously identified seventeen ($\approx 27^{+6}_{-9}$ percent) sources to be hosting AGN activity in galaxies with $L_{\text{IR}} > 3 \times 10^8 L_\odot$, using [NeV] $\lambda 14.32$ μm emission as a robust AGN indicator. Using the SMBH mass and AGN bolometric luminosity estimates derived here, we discuss the relative mass accretion rates of these seventeen AGNs and use them to derive the average present-day growth times of SMBHs in the very nearby Universe. Furthermore, we evaluate the unique contribution that our new optically unidentified AGNs make to the space density of active SMBHs in the local Universe, which have until now been previously derived from large-scale optical surveys (e.g., H04; Greene & Ho 2007).

5.1 Derived AGN Properties and Relative Mass Accretion Rates

In Fig. 3 we plot $L_{\text{bol,AGN}}$ against our adopted $M_{\text{BH}}$ estimates (the associated 1-σ errors for $M_{\text{BH}}$ measurements are described in section 3) for the 17 AGNs in our volume-limited sample. $L_{\text{bol,AGN}}$ is inferred from either accurate intrinsic high-quality hard X-ray (2-10 keV) constraints (where available) or AGN-produced [OIV] $\lambda 25.80$ μm emission (see §4.1 and 4.2). The $L_{\text{bol,AGN}}$ 1-σ errors for the sources with hard X-ray constraints are the result of combining the uncertainty in the $L_{\text{X,2-10keV}}$ measurement with that of the mean spread in the bolometric correction factor employed from Marconi et al. (2004). For those AGNs with $L_{\text{bol,AGN}}$ derived from $L_{\text{OIV}},$ the error is derived from the uncertainty in $L_{\text{OIV}}$ as quoted in GA09 combined in quadrature with the intrinsic scatter of the empirical [OIV]–$L_{\text{bol,AGN}}$ relation (equation 4).

Seven objects within the sample have been classified as optical AGNs from previous surveys (see GA09 and Table 1) using typical optical emission-line diagnostics (e.g., the Baldwin-Phillips-Terlevich diagnostic diagrams; Baldwin et al. 1981); however, all have detected [NeV] $\lambda 14.32$ μm emission, and thus are unambiguously identified to host AGNs at mid-IR wavelengths (GA09). We find that with the exception of NGC 5128 (Centaurus A), our sample is dominated by AGNs with SMBHs in the mass range $M_{\text{BH}} \approx (0.1–5) \times 10^7 M_\odot$ (median of $M_{\text{BH}} \approx 7 \times 10^6 M_\odot$). Due to the irregular structure of one of the galaxies in the sample (NGC 5195), $M_{\text{BH}}$ is poorly determined; in Fig. 3 we plot $M_{\text{BH}}$ estimates from both the M-σ, and $M_{\text{BH}}$–$L_{\text{K,bol}}$ relations (connected blue-dashed line).

We find the AGNs in our sample are spread over a wide range of bolometric luminosities, $L_{\text{bol,AGN}} \approx 10^{40}$–$10^{45}$ erg s$^{-1}$. To assess the relative mass-accretion rates of the sample ($L_{\text{bol,AGN}}/L_{\text{Edd}} \sim \eta$), we over-plot lines of constant Eddington ratios ($\eta \approx 10^{-2}, 10^{-3}, 1.0$; derived following Reeves 1982) and their associated mass-doubling times ($t \approx 30, 0.3, 0.03$ Gyrs, respectively). Given the large range

![Figure 4. $M_{\text{BH}}$ is plotted against the characteristic mean mass doubling time ($t_{\text{MD}}$) of a SMBH in units of Hubble-time ($t_H$) for the $D < 15$ Mpc AGNs. Growth time errors are calculated from the log-normal standard deviations of the sample. For comparison, the growth time function of H04 is shown (dashed line). We find good agreement with H04 over the region $M_{\text{BH}} \approx 0.3–10 \times 10^8 M_\odot$ and further extend the growth time constraints to lower SMBH masses ($M_{\text{BH}} \approx 10^6 M_\odot$).](image)

in bolometric luminosities, it is not surprising that the AGNs in the sample are found to be accreting at rates covering over 5 orders of magnitude ($\eta \approx 10^{-5}–1$). With the exception of a few AGNs, the observed range in Eddington ratios is found to be roughly consistent with those found by H04 for active galaxies (solid contours in Fig. 4). As our work is not limited by the spectral resolution of the SDSS (i.e., with a limit of $M_{\text{BH}} \geq 3 \times 10^6 M_\odot$), we show in Fig. 3 that significant accretion, $\eta \approx 10^{-3}$ (i.e., radiatively efficient accretion systems; e.g., thin discs) occurs onto SMBHs with $M_{\text{BH}} \approx (1–3) \times 10^6 M_\odot$. The majority of these low-mass, rapidly-accreting SMBHs are hosted in late-type, disc-dominated spiral galaxies (Sc–Sd). By contrast, it is generally assumed that gas-rich late-type spirals are preferentially inactive galaxies and that a large bulge may be a necessary component for the existence of a SMBH, and thus a luminous AGN. Furthermore, of the four AGNs within the sample with SMBHs consistent with $M_{\text{BH}} \approx 10^6 M_\odot$, we find that three sources are not identified as AGNs in sensitive optical surveys. This indicates that significant SMBH accretion may be missed by statistically-large optical surveys such as H04 even if the spectral resolution was sufficient to identify SMBHs down to $M_{\text{BH}} \approx 10^6 M_\odot$.

For the subset of our AGN sample which host SMBHs with $M_{\text{BH}} \geq 3 \times 10^6 M_\odot$, we find that many of the optically unidentified AGNs are accreting at relatively low Eddington ratios ($\eta \approx 10^{-3}$), and are unlikely to make a significant additional contribution to the present-day growth of SMBHs. However, these same AGNs may form part of a separate, underlying population of radiatively inefficient accretion sys-
tems such as advection dominated accretion flows (ADAFs; e.g., Narayan & Yi 1994) or those which contain optically-thick slim-discs. Further spectral analysis of the X-ray data may distinguish between these particular accretion systems, but is beyond the scope of these analyses (see Goulding et al. in preparation).

5.2 The Present-Day Growth of SMBHs

Using the relative mass accretion rates estimated for our sample (Fig. 3), we can infer the volume-averaged growth time of SMBHs in the local Universe. Assuming a mean Kerr spin parameter ($a$) for our sample of $a \approx 0.67$ (e.g., Treister & Urry 2006; Hopkins et al. 2007), i.e., an accretion efficiency ($\epsilon$) of $\approx 0.1$, the characteristic mass doubling time ($t_{2M}$) of a SMBH accreting matter at the Eddington limit is $t_{2M} \approx 30$ Myrs (Rees 1984). Under the further assumption that $a$, and hence $\epsilon$, does not vary significantly for changes in $M_{\text{BH}}$ (King et al. 2008), we assess the present-day growth rate of SMBHs. Following H04, we calculate and extend to lower masses ($M_{\text{BH}} < 3 \times 10^6 M_\odot$) the integrated growth of SMBHs. Growth time errors are calculated from the log-normal standard deviations of the sample. We note here that we also include the optically unidentified AGNs which would not be detected in the SDSS.

In Fig. 3 we find that the mean growth time for low-mass SMBHs ($M_{\text{BH}} \approx 10^6 M_\odot$) is $\approx 6^{+6}_{-3}$ Gyr, which is consistent with these AGNs growing on time-scales similar to that of the age of the Universe. Our results are found to be broadly consistent with a simple extrapolation of the growth times calculated by H04 to $M_{\text{BH}} \approx 10^6 M_\odot$ (dashed line in Fig. 3). Thus, the AGNs hosting SMBHs in the mass range $M_{\text{BH}} \approx 10^6 - 10^9 M_\odot$, which are dominated by optically unidentified AGNs (see Fig. 3), are acquiring a significant proportion of their mass in the present-day, and are amongst the most rapidly growing SMBHs in the local Universe. Furthermore, we find our derived growth times of SMBHs with $M_{\text{BH}} \gtrsim 3 \times 10^6 M_\odot$ are in good agreement with those presented in H04, with mean growth times of $t_{2M} \approx 47^{+29}_{-18}$ and $\approx 198^{+90}_{-43}$ Gyr for AGNs with $M_{\text{BH}} \approx (0.3 - 3)$ and $(3 - 30) \times 10^8 M_\odot$, respectively.

5.3 Space-Density of AGNs in the local Universe

An accurate active SMBH mass function, especially for lower mass SMBHs ($M_{\text{BH}} \approx 10^6 M_\odot$), is crucial for extending our understanding of the role played by accretion in the growth of SMBHs across cosmic time. In this section we calculate the space density of active SMBHs for our sample and compare it to complimentary optical studies of local narrow-line (NL; H04) and broad-line (BL; Greene & Ho 2007) AGNs.

The derived volume-weighted space density for our active SMBHs (filled squares), which is dominated by NL-AGNs, is found to be significantly greater (a factor of $\approx 100$) than the SMBH density of BL-AGNs (filled circles) presented in Greene & Ho (2007) in the mass region $M_{\text{BH}} \approx (0.9 - 0.90) \times 10^6 M_\odot$. The significant increase in active SMBH density when compared to the BL-AGN density of Greene & Ho (2007) is partially to be expected due to the greater relative sensitivity of our Spitzer-IRS observations coupled with the greater abundance of observed Seyfert 2 to Seyfert 1 galaxies identified in the local Universe. However, we may be rejecting broad-line objects, and thus finding a lower Seyfert 1:Seyfert 2 ratio than is representative in the local Universe.

As we note in §2, given the luminosity limit imposed in our volume-limited survey ($L_{\text{IR}} \gtrsim 3 \times 10^9 L_\odot$), we do not include the dwarf like systems which are likely to host the very smallest SMBHs ($M_{\text{BH}} < 10^5 M_\odot$). With an adjustment for the distance model adopted in this work, an examination of the Palomar survey (Ho et al. 1997a) shows there are possibly three optical Seyferts (NGC 185; NGC 1058; NGC 4395) with $M_{\text{BH}} < 10^5 M_\odot$ which are not included in our sample due to our lower luminosity limit. However, we also note that NGC 185 has since been re-classified as an HII galaxy (Ho & Ulvestad 2001).

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In comparison to the active SMBH mass function containing the optically identified NL-AGNs of H04 (dotted line), we also find a significantly larger space-density of SMBHs. We find that the space-density of active SMBHs identified in the mid-IR is roughly constant in the mass range $M_{\text{BH}} \approx (0.9 - 0.90) \times 10^6 M_\odot$ with a value of $\Phi \approx 6.3 \times 10^{-4} \text{Mpc}^{-3} \log M_{\text{BH}}$. This space-density of AGNs is a factor of $\approx 10$ greater than that estimated by H04 over the same $M_{\text{BH}}$ range. Since we find only two of the 17 AGNs in our sample are sufficiently luminous/unobscured to be detected in the SDSS survey, we determine that this is consistent with our results. We further suggest that the space density derived here may still be a lower-limit for the number of NL-AGNs in the local Universe. A further examination of the (distance-model adjusted) Palomar survey suggests that at least four further NL-AGNs are not included in our volume-limited survey. Of these, two (NGC 3486; IRAS 6048+2258) and the total mass function of local SMBHs by Marconi et al. (2004).

Following Greene & Ho (2007), in the top panel of Fig. 3 we plot the volume-weighted space density, $\Phi$ against $M_{\text{BH}}$ in mass bins of 0.5 dex. The volume, $V$, encompassed by the GA09 sample to $D < 15$ Mpc is $V \approx 1.3 \times 10^8 \text{Mpc}^3$. As we note in §2, given the luminosity limit imposed in our volume-limited survey ($L_{\text{IR}} \gtrsim 3 \times 10^9 L_\odot$), we do not include the dwarf like systems which are likely to host the very smallest SMBHs ($M_{\text{BH}} < 10^5 M_\odot$). With an adjustment for the distance model adopted in this work, an examination of the Palomar survey (Ho et al. 1997a) shows there are possibly three optical Seyferts (NGC 185; NGC 1058; NGC 4395) with $M_{\text{BH}} < 10^5 M_\odot$ which are not included in our sample due to our lower luminosity limit. However, we also note that NGC 185 has since been re-classified as an HII galaxy (Ho & Ulvestad 2001).
NGC 4565 lack high-resolution Spitzer-IRS spectroscopy of the central region (as noted in Table 2 of GA09), one (NGC 3031; L_{IR} \approx 2.8 \times 10^9 \ L_\odot) lies fractionally below our luminosity limit for this survey, and NGC 4258 is not included in the RBGS due to its extremely large angular size.

Whilst the main focus of this paper is to compare SMBH statistics derived from mid-IR and optical detection techniques, it is prudent to note that the majority of AGN space densities calculated at higher redshifts are typically derived from sources detected in wide-field X-ray surveys. Hence, we now establish whether our mid-IR active space density may be missing a significant fraction of X-ray detected AGNs. Recently, a comparison between high-resolution Spitzer-IRS spectroscopy and X-ray detected AGNs, was made by Dudik et al. (2009) for a large sample of optically classified LINERS. Dudik et al. (2009) reported inconsistencies between [NeV] non-detections and the presence of hard X-ray nuclear emission in a subset of their sample. However, they conclude that the limited sensitivity of their mid-IR observations may be driving their observed result. Using the relations between high-ionisation mid-IR emission and hard X-ray luminosity (M08; GA09; Goulding et al. in
preparation), we suggest that a detected [NeV] luminosity of $L_{[\text{NeV}]} \approx 10^{38}$ erg s$^{-1}$ in an AGN (i.e., the limiting luminosity in Dudik et al. 2009) would be equivalent to a hard X-ray luminosity of $L_{X,2–10keV} \approx 5 \times 10^{40}$ erg s$^{-1}$; indeed, almost all of the X-ray detected AGNs which lack significant [NeV] emission in Dudik et al. (2009) are below this threshold. We thus conclude that with sufficiently sensitive high-resolution mid-IR spectroscopy, there is currently no conclusive evidence to suggest that X-ray detected Type-2 AGNs lack significant [NeV] emission in the mid-IR. Hence, with the exception of some Seyfert 1 galaxies (as discussed previously), it is unlikely that our derived mid-IR space density lacks significant numbers of X-ray detected AGNs with [NeV] emission below our sensitivity limit.

Given our comparatively small volume to that considered by using the SDSS (e.g., Greene & Ho 2007; H04), we further validate our derived space density of active SMBHs by robustly testing our results to find if: 1) our sample is over-dense, and thus strongly subject to cosmic variance; or 2) given the modest errors associated with our $M_{\text{BH}}$ estimates, the derived space-density is strongly subject to scattering of objects in our defined binning structure. We discuss these analyses in the Appendix. Briefly, we find that our sample is broadly representative of galaxies to $z \approx 0.3$, and further show that even in our most pessimistic case, we find an increase in our derived space-density of at least a factor of $\approx 2$ (maximum increase by a factor of $\approx 11$) at $M_{\text{BH}} \approx 3 \times 10^6 M_\odot$ over the optical NL space-density of H04.

5.4 The volume-weighted local AGN fraction

The ratio of the space densities of the active SMBH to total SMBH mass function (i.e., the volume-weighted local active SMBH fraction) is shown in the lower panel of Fig. 5. We calculate an overall active SMBH fraction of $\approx 25^{+15}_{-29}$ percent for SMBHs of $M_{\text{BH}} \approx (0.5–500) \times 10^6 M_\odot$ down to our [NeV] completeness limit ($L_{[\text{NeV}]} \approx 10^{38}$ erg s$^{-1}$). We find that this fraction is consistent with being constant throughout this $M_{\text{BH}}$ range. However, given our detection sensitivity limit, we are unable to probe lower Eddington ratios for AGNs hosting smaller SMBHs. Instead, we consider the effect of the AGN fraction for a fixed value of Eddington ratio (e.g., $\eta > 10^{-3}$; i.e., thin-disc accretion systems) and find tentative evidence that the AGN fraction ($\approx 10^{+8}_{-6}$ and $\approx 8^{+10}_{-5}$ percent) may increase with decreasing SMBH mass ($M_{\text{BH}} \approx 10^6–10^7 M_\odot$ and $M_{\text{BH}} \approx 10^7–10^8 M_\odot$ bins, respectively).

For the lowest-mass SMBHs ($M_{\text{BH}} \approx (5–30) \times 10^5 M_\odot$), we estimate an overall non-negligible volume-weighted AGN fraction of $18^{+7}_{-2}$ percent, potentially showing that a considerable proportion of small bulge (and pseudo-bulge) galaxies (i.e., late-type spiral galaxies; Sc-Sd) host AGN activity. It has been previously suggested by H04 and Greene & Ho (2007) that the AGN fraction may peak at $M_{\text{BH}} \approx (0.7–2) \times 10^7 M_\odot$. However, with the inclusion of the additional low-mass optically unidentified AGNs (see Fig. 3), we find that the AGN fractions are consistent with remaining constant or even increasing for $M_{\text{BH}} < 10^7 M_\odot$.

As noted in §5.3, the space-density of AGNs, and hence the local AGN fraction derived in this work, may only be a lower-limit given the nature of our volume-limited survey which by definition does not include IR-faint systems. To improve upon these current source statistics, a larger sample of late-type spiral galaxies would be required to investigate our findings further. With the greater sensitivity and resolving power proposed for the next generation of space-based mid-IR spectrographs, for example the Space Infrared Telescope for Cosmology and Astrophysics (SPICA) and the mid-IR instrument (MIRI) on-board the James Webb Space Telescope (JWST) surveys such as these can be continued and extended to study more distant (i.e., greater volumes) and heavily obscured AGNs.

6 CONCLUSIONS

We have presented the mean growth times and volume-weighted space density of active SMBHs in the local Universe. Our sample of 17 AGNs was derived from a sensitive volume-limited mid-IR spectral survey of all IR bright galaxies to $D < 15$ Mpc carried out using the NASA Spitzer Space Telescope (see GA09 for further details on the sample selection). The most accurate SMBH masses available for the objects are compiled from a variety of sources. For the three AGNs without published $M_{\text{BH}}$ estimates, we use a bulge/disc decomposition method to determine the bulge luminosity and hence a SMBH mass (see section 3). Our main findings are the following:

1. Using combined mid-IR emission line and high-quality hard X-ray constraints, we have derived accurate measurements of the intrinsic luminosities of our sample of AGNs (see section 4). In conjunction with the well-established SMBH measurements from previous studies and our own estimates from our aforementioned bulge/disc decomposition method, we have assessed the relative mass accretion rates of our sample. Due to our high sensitivity and the ability to probe low SMBH masses, we find that significant mass accretion ($\eta > 10^{-3}$) occurs onto SMBHs with $M_{\text{BH}} \approx 10^6 M_\odot$, the majority of which would not be detected in even the most sensitive optical surveys. See section 5.1.

2. Using our derived relative mass accretion rates for the sample and assuming a typical accretion efficiency of $\epsilon \approx 0.1$, we assessed the characteristic mean mass doubling times ($t_{2M}$) for AGNs in the very nearby Universe. For AGNs hosting SMBHs with $M_{\text{BH}} \approx (0.5–50) \times 10^7 M_\odot$ we find consistent growth times ($t_{2M} \approx 47–198$ Gyr) with those of the NL-AGNs identified in the SDSS (H04). However, we also find that SMBHs with $M_{\text{BH}} \approx 5 \times 10^6 M_\odot$ (i.e., below the spectral resolution limit of the SDSS), are amongst the most rapidly growing SMBHs in the local Universe, with present-day growth times consistent with (and possibly less than) the current age of the Universe ($t_{2M} \approx 6^{+6}_{-3}$ Gyr). See section 5.2.

3. To assess the incidence of this population of low mass, rapidly growing SMBHs, we constructed a local space density function of active SMBHs. We find that active SMBHs may be at least a factor $\approx 2$ more common than previously identified in NL-AGN surveys using SDSS data. Furthermore, we estimate a non-negligible space density

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12 See http://www.ir.isas.jaxa.jp/SPICA/
13 See http://www.roe.ac.uk/uktac/consortium/miri/
for low mass SMBHs ($M_{\text{BH}} \approx 10^6 M_\odot$) of $\Phi \approx 6 \times 10^{-4} Mpc^{-3} \log M_{\text{BH}}^{-1}$ which is consistent with the space density of more massive active SMBHs ($M_{\text{BH}} \approx 10^7 M_\odot$; i.e., those previously determined to be the most rapidly accreting population of SMBHs). See section 5.3

(4) Using a local total SMBH mass function (Marconi et al. 2004), we estimate a mean volume-weighted local AGN fraction of $\approx 23^{+22}_{-11}$ percent, which remains relatively constant in the mass range $M_{\text{BH}} \approx (1-10) \times 10^8 M_\odot$. However, when only considering the SMBHs with $\eta > 10^{-3}$ (i.e., radiatively efficient accretion systems), we find tentative evidence for an increasing AGN fraction ($\approx 16^{+9}_{-7}$ and $\approx 8^{+10}_{-6}$ percent) with decreasing SMBH mass (for $M_{\text{BH}} \approx 10^6-10^7 M_\odot$ and $M_{\text{BH}} \approx 10^7-10^8 M_\odot$, respectively). See section 5.4

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APPENDIX A: VALIDATION OF DERIVED SPACE-DENSITY OF ACTIVE SMBHS

As the space-density of active SMBHs derived in this work is significantly larger than that found in previous studies, in this section we attempt to validate our results by discussing possible limitations and additional sources of error which may exist in these analyses: 1) whilst the sensitivity of the data used in this survey is high, the volume considered is relatively small compared to that of the SDSS, thus our results may be subject to cosmic variance; 2) given the modest errors associated with the $M_{\text{BH}}$ estimates, the adopted $M_{\text{BH}}$ binning structure is likely to be subjective and thus degenerate towards objects scattering between the defined bins.

Under these assumptions, in §A1 we investigate whether the sample is indeed representative of the local Universe and in §A2 we discuss the construction of a Monte Carlo simulation to assess the effect of our adopted $M_{\text{BH}}$ binning.

A1 Is our sample representative of the local Universe?

Given the large incidence of AGNs within our sample it is possible that the volume considered in our sample is over-dense compared to other regions in the local Universe. The construction of the original sample of 64 bolometrically luminous galaxies in GA09 was designed to be complete down to the flux-limit of the Revised Bright Galaxy Sample. This imposed a distance constraint of $D < 15$ Mpc (see Fig. 1 of GA09), and hence did not include the Virgo cluster at $D \approx 16$ Mpc, and thus our sample does not incorporate known local over-densities; however, this volume may not be representative of the Universe at large.

To robustly test our considered volume ($V \approx 1.3 \times 10^4$ Mpc$^3$), we constructed a total SMBH space density function for all galaxies to $D < 15$ Mpc and compared this to the total local (active+inactive) SMBH mass function of Marconi et al. (2004) derived from the luminosity function of local galaxies. Given the large co-moving volume ($V_c \approx 1000$ Gpc$^3$) considered in Marconi et al. (2004), their derived SMBH mass function is unlikely to suffer from significant cosmic variance.

Our total SMBH mass function was formulated using all galaxies identified in the NASA/IPAC Extra-galactic Database (NED) to $D < 15$ Mpc with a total K-band luminosity of $L_{K,gal} \gtrsim 1.5 \times 10^9 L_\odot$. The luminosity threshold is designed to include all galaxies which could potentially host a SMBH with $M_{\text{BH}} \gtrsim 10^8 M_\odot$ using the $M_{\text{BH}}-L_{K,gal}$ relation. This conservative lower limit assumes that $L_{K,gal} = L_{K,bul}$ (i.e., that all galaxies in the sample have an early-type galaxy classification). In reality, the majority of the sources identified in NED to $D < 15$ Mpc are late-type galaxies (i.e., $L_{K,gal} \gg L_{K,bul}$), and therefore the $M_{\text{BH}}$ limit is likely to be $M_{\text{BH}} \lesssim 10^7 M_\odot$. To $D < 15$ Mpc, we identify 105 galaxies which potentially host a SMBH with $M_{\text{BH}} \gtrsim 10^8 M_\odot$. To estimate $M_{\text{BH}}$ in each of these galaxies, we use the associated Hubble-type from the Third Reference Catalogue of Bright Galaxies (de Vaucouleurs et al. 1991) to a mean bulge/disc ratio (e.g., Benson et al. 2007) and establish an individual bulge luminosity based on galaxy-type and $L_{K,bul}$. We convert the estimated bulge luminosity to $M_{\text{BH}}$ using the $M_{\text{BH}}-L_{K,bul}$ relation and construct a total SMBH mass function. Whilst this rather crude estimation carries large associated errors, we still find very good agreement (a mean variance of 0.1 dex) with the SMBH mass function of Marconi et al. (2004) throughout the mass range $M_{\text{BH}} \approx 10^6-10^9 M_\odot$. Thus, to first-order, our volume-limited sample does not appear to be over-dense and/or subject to strong cosmic variance, and hence, is broadly representative of the typical field-galaxy population in the local Universe.

We further validate this conclusion by estimating the active SMBH mass function for those AGNs which are sufficiently optically bright to be included in H04. Under this assumption, the derived space densities of active SMBHs for this sample and that of H04 should be comparable if our sample is indeed representative of a field-galaxy population. We find that using the H04 detection limit inferred from Fig. 4 our new estimated space density is consistent with H04 ($\Phi \approx 10^{-4}$) throughout the mass range considered here ($M_{\text{BH}} \approx (0.05-30) \times 10^8 M_\odot$), although there are still considerable uncertainties given the small number statistics inherent to our sample.

A2 Monte Carlo Analysis of $M_{\text{BH}}$ binning

We have established that our sample and the estimate of the SMBH mass density (both from active and inactive galaxies) do not appear to be subject to over-densities caused by cosmic variance. We therefore now investigate the effect of small-number statistics on our results which are inherent in relatively small samples such as this.

We show in §3 that accurate estimates of $M_{\text{BH}}$ for our sample of AGNs from a homogeneous method is difficult to achieve. Thus, our adopted values of $M_{\text{BH}}$ are derived from a variety of methodologies. In Fig. 3 the AGNs in the sample are placed into equal bins of $M_{\text{BH}}$ with width 0.5 dex. However, this binning process does not allow for the error inherent to each individual $M_{\text{BH}}$ measurement. Hence, some objects may scatter out of one defined bin and into another. The bin width we employ in our relatively modest sample (64 objects) may therefore be subjective and requires testing. Here we use a Monte Carlo analysis to assess the effect of the scattering of AGNs into different $M_{\text{BH}}$ bins on the derived space density of active SMBHs in the local Universe ($\Phi$).

Our Monte Carlo analysis calculates $\Phi$ by selecting a random set of SMBH masses from gaussian probability distributions constructed using our adopted $M_{\text{BH}}$ masses and their associated 1-σ errors given in column 12 of Table 1. The binning structure was designed to incorporate at least one object in each bin from 250,000 realisations of the simulation across the considered mass range of the sample ($M_{\text{BH}} \approx (0.5-500) \times 10^8 M_\odot$). We impose the upper mass limit as we have only one high-mass AGN (NGC 5128; $M_{\text{BH}} \approx 2.4 \times 10^9 M_\odot$) in the sample. Thus, we determine that given our sample distribution the maximum number of equal width bins to be nine (i.e., equal bin sizes of 0.33 dex with $\geq 1$ object). Hence, we use our simulation to conservatively assess the maximum error on our calculated space density of active SMBHs.

The error shown in Fig. 5 (shaded region) is the standard deviation of the 250,000 simulations in our volume-limited sample combined in quadrature with the Poisson error determined from the number counts in our real sample. We found 250,000 realisations to be sufficient, since at this
level the maximum variation in the standard deviation over multiple runs was less than $10^{-6}$.

We find that the mean spread in the derived value of $\Phi$ from our volume-limited sample is $\approx 0.85$ dex, and as predicted appears to be subject to some scattering of $M_{BH}$. However, using our conservative error analyses, we show that the space density of active SMBHs found from our sample is consistently greater than that found for optical narrow-line AGNs (H04) with $M_{BH} < 10^8 M_\odot$. Thus, to first-order, we find a significant increase (of a minimum factor of $\approx 2$ and a maximum of $\approx 80$) in the space density of active SMBHs in the local Universe in the mass range $M_{BH} \approx (0.5-100) \times 10^6 M_\odot$. For $M_{BH} > 10^8 M_\odot$ our errors increase significantly due to very limited source statistics (see §4.1). We thus conclude that while this survey is subject to the scattering of objects through the $M_{BH}$ bins, even in our most pessimistic case, we still find a significant increase in the space density of active SMBHs in the local Universe compared to that found in large-scale optical NL AGN studies.