Analytical Predictions for Statistics of Cosmic Shear: Tests Against Simulations

Patrick Valageas\textsuperscript{1}, Andrew J. Barber\textsuperscript{2}, Dipak Munshi\textsuperscript{3,4}
\textsuperscript{1} Service de Physique Théorique, CEA Saclay, 91191 Gif-sur-Yvette, France
\textsuperscript{2} Astronomy Centre, University of Sussex, Falmer, Brighton, BN1 9QJ, United Kingdom
\textsuperscript{3} Institute of Astronomy, Madingley Road, Cambridge, CB3 OHA, United Kingdom
\textsuperscript{4} Astrophysics Group, Cavendish Laboratory, Madingley Road, Cambridge CB3 OHE, United Kingdom

20 March 2022

\textbf{ABSTRACT}

Weak gravitational lensing surveys are rapidly becoming important tools to probe directly the mass density fluctuations in the universe and its background dynamics. Earlier studies have shown that it is possible to model the statistics of the convergence field on small angular scales by suitably modeling the statistics of the underlying density field in the highly non-linear regime. We extend such methods to model the complete probability distribution function, PDF, of the shear as a function of smoothing angle. Our model relies on a simple hierarchical \textit{Ansatz} for the behaviour of the higher-order correlations in the density field. We compare our predictions with the results of numerical simulations and find excellent agreement for different cosmological scenarios. Our method provides a new way to study the evolution of non-Gaussianity in gravitational clustering and should help to break the degeneracies in parameter estimation based on analysis of the power spectrum alone.

\textbf{Key words:} Cosmology: theory – gravitational lensing – large-scale structure of Universe Methods: analytical – Methods: statistical – Methods: numerical

\section{INTRODUCTION}

Recent detection of weak lensing (see, e.g., Bartelmann & Schneider, 2001, for a review) due to the large-scale structure in the universe has presented a new opportunity to study cosmology (see, e.g., Bacon, Refregier & Ellis, 2000, Hoekstra et al., 2002, Van Waerbeke et al., 2000, and Van Waerbeke et al., 2002). With a tight control on various systematics these studies were able to relate the weak distortions in the images of high-redshift galaxies with the intervening large-scale mass density inhomogeneities in the universe. Such studies which are statistical in nature (as opposed to strong lensing studies) provide us with a unique way to probe the statistical properties of the underlying density distribution. Traditionally, the study of gravitational clustering in the quasi-linear and non-linear regimes was done by analyzing galaxy catalogues and were plagued by the question of bias associated with the galaxy distribution. Now, therefore, one needs a prescription for galaxy bias to relate the galaxy distribution to the underlying dark matter distribution. Gravitational lensing allows us to bypass this problem as it directly probes the mass distribution. Pioneering works in this direction were done by Blandford et al. (1991), Miralda-Escudé (1991), and Kaiser (1992), based on the earlier work by Gunn (1967).

A key tool for the understanding of cosmological weak lensing has been provided by the development of numerical techniques to follow null geodesics in a perturbed model universe. Such numerical studies typically employ $N$-body simulations through which ray-tracing experiments are conducted (see, e.g., Schneider & Weiss, 1988, Jarozyński et al., 1990, Lee & Paczynski, 1990, Jarożynski, 1991, Babul & Lee, 1991, and Blandford et al., 1991). Building on the earlier work of Wambsganss et al. (1995 and 1997), detailed numerical studies were done by Wambsganss, Cen & Ostriker (1998). Other recent studies using ray-tracing have been conducted by Premadi, Martel & Matzner (1998), Van Waerbeke, Bernard & Mellier (1999) and White & Hu (2000). Recently, Couchman, Barber & Thomas (1999) have developed a new technique based on computing the full three-dimensional shear at locations along the lines of sight. Their method combines the computed shear matrices to produce the required lensing statistics for sources at the selected redshift. The method is complimentary to ray-tracing and has been fully implemented. Excellent agreement between the results obtained using Couchman et al.’s (1999) method and
analytical predictions have been reported by Barber (2002) and Barber & Taylor (2002).

On the other hand, analytical approaches have also been developed to elucidate the link between the statistical distortion of distant sources by weak lensing and the properties of the large-scale structure of the universe. Most of these analytical works have focussed on large smoothing angles where perturbative calculations are valid (e.g., Villumsen, 1996, Bernardeau et al., 1997, Jain & Seljak, 1997, Kaiser, 1998, Van Waerbeke, Bernardeau & Mellier, 1999, and Schneider et al., 1998). Small angular scales are more difficult to handle since they probe the non-linear regime where the density field can no longer be described through perturbative expansions which start behaving badly (e.g., Valageas, 2002b).

Thus, for ongoing surveys with a small sky coverage perturbative results are unlikely to be sufficient. Therefore, we need theoretical predictions at small angular scales where the signal will dominate the noise.

Unfortunately, we lack a complete understanding of gravitational clustering in the highly non-linear regime since it has not been possible so far to solve the Vlasov-Poisson system which describes the collisionless gravitational dynamics. Nevertheless, several phenomenological models have been developed to describe the non-linear density field. An interesting family of such models, often called “hierarchical models”, assumes a tree hierarchy for many-body correlation functions. Then, each model is specified by the way it assigns weights to trees of the same order but of different topologies (Fry, 1984, Schäffer, 1984, Bernardeau & Schäffer, 1992, and Szapudi & Szalay, 1993 and 1997). Note that the evolution of the two-point correlation function in all such approximations is left arbitrary. On the other hand, independent studies by various authors (Hamilton et al., 1991, and Peacock & Dodds, 1994 and 1996) suggest an accurate fitting formula for the evolution of the two-point correlation function based on simple non-local scaling arguments. Combining a hierarchical Ansatz with such a non-local scaling for the two-point correlation one obtains a complete statistical description of gravitational clustering in the highly non-linear regime (e.g., Valageas & Schäffer, 1997, Munshi et al., 1999a, Munshi, Coles & Melott, 1999b,c, and Munshi, Melott & Coles 1999d).

Such an approach was initiated by Hui (1999) for the study of the three point correlation function for convergence maps and was later extended to compute the full probability distribution function (Valageas, 2000a,b, and Munshi & Jain, 2000, 2001), the associated bias (Munshi, 2000) and the cumulant correlators associated with such distributions (Munshi & Jain, 2000). See also Munshi & Wang (2003) for the extension of these studies to cosmological scenarios with dark energy. Next, such studies were extended to handle the compensated filter appropriate for the aperture-mass, $M_{ap}$, which is more useful in weak lensing surveys (Bernardeau & Valageas, 2000). Error estimations were also performed based on such a hierarchical Ansatz to optimize survey strategies (Munshi & Coles, 2003), while implications for Type Ia Supernova studies were carried out by Valageas (2000a) and Wang, Holz & Munshi (2002).

So far most analytical calculations have focussed on convergence maps where analytical results are simplified due to the fact that the convergence is a scalar. Similar progress in the case of the shear was lacking due to the spinorial nature of shear fields. However from an observational point of view, shear maps are a direct outcome of surveys whereas map-making for the convergence field can be problematic due to the non-trivial topology of the survey area. Therefore, extending previous studies of weak gravitational lensing based on hierarchical models, we introduce in this article a model, which we also call a “stellar model,” to evaluate the probability distribution function, PDF, of the shear. We test our analytical predictions against numerical simulations and we show that our simple model provides a good description of the shear PDF at small angular scales. Our approach is complimentary to studies related to lower order statistics of the shear field (Takada & Jain, 2003, Zaldarriaga & Scoccimarro, 2003, Schneider & Lombardi, 2003, and Bernardeau, Mellier & van Waerbeke, 2002) which also aim to detect non-Gaussianity from ongoing surveys.

This paper is organized as follows. Section 2 briefly describes the weak gravitational lensing effects as measured through the convergence, $\kappa$, and the shear, $\gamma$. In Section 3, we recall the relationship between PDFs and cumulants (which are the basic tools of our analytical method) and we present our model for the many-body density correlations. This allows us to compute in Section 4 the PDFs for the shear components and for the shear modulus. Section 5 contains a brief explanation of the method used to obtain the lensing statistics from the numerical simulations. In Section 6, we make the comparison between these numerical results and our analytical predictions, and we conclude in Section 7 by discussing our results and future work.

2 DISTORTIONS INDUCED BY WEAK GRAVITATIONAL LENSING

2.1 Shear tensor

The gravitational lensing effects produced by density fluctuations along the trajectory of a photon lead to an apparent displacement of the source and to a distortion of the image. Thus, light coming from a direction $\theta$ is deflected by a small angle $d\theta$. However, the observable quantities are not the displacements $d\theta$ themselves but the distortions induced by these deflections, which are given by the symmetric shear matrix (e.g., Jain, Seljak & White, 2000):

$$\Phi_{ij} = \frac{\partial \theta_i}{\partial \theta_j} = -2 \int_0^{\chi_s} d\chi \frac{D(\chi) D(\chi_s - \chi)}{D(\chi_s)} \nabla_i \nabla_j \phi(\chi).$$

Here $\chi$ is the radial comoving coordinate (and $\chi_s$ corresponds to the redshift, $z_s$, of the source):

$$d\chi = \frac{dz}{\sqrt{\Omega_\Lambda + (1 - \Omega_m - \Omega_\Lambda)(1 + z)^2 + \Omega_m(1 + z)^3}}$$

while the angular distance $D$ is defined by:

$$D(\chi) = \frac{\pi m}{\sin_k \left( \left| 1 - \Omega_m - \Omega_\Lambda \right|^{1/2} \frac{H_0}{c} \chi / c \right)} \left| 1 - \Omega_m - \Omega_\Lambda \right|^{1/2}.$$

where $\sin_k$ means the hyperbolic sine, sinh, if $(1 - \Omega_m - \Omega_\Lambda) > 0$, or sine if $(1 - \Omega_m - \Omega_\Lambda) < 0$; or $(1 - \Omega_m - \Omega_\Lambda) = 0$, then $D(\chi) = \chi$. The gravitational potential, $\phi$, is related to the fluctuations of the density contrast, $\delta$, by Poisson’s equation:

$\odot 0000 RAS, MNRAS 000, 000–000$
Following Valageas (2000a, b) it is convenient to define the contrast along the line of sight. Then, we can see from eq.(8), 
\[ z_\delta = \frac{\rho(x) - \rho}{\rho} \]
where \( \rho \) is the mean density of the universe. In eq.(1) we used the weak lensing approximation; the derivatives \( \nabla_i \nabla_j \phi(\chi) \) of the gravitational potential are computed along the unperturbed trajectory of the photon. This assumes that the components of the shear tensor are small but the density fluctuations \( \delta \) can be large (Kaiser, 1992). The shear tensor \( \Phi_{i,j} \) is usually decomposed into its trace, \( \kappa \), and the shear components \( \gamma_1, \gamma_2 \), defined by:
\[ \kappa = -\frac{\Phi_{1,1} + \Phi_{2,2}}{2} \]
and
\[ \gamma_1 = -\frac{\Phi_{1,1} - \Phi_{2,2}}{2}, \quad \gamma_2 = -\Phi_{1,2} \quad \gamma = \gamma_1 + i \gamma_2. \]
Note that since the shear components, \( \gamma_1 \) and \( \gamma_2 \), and their signs can be exchanged through appropriate rotations of the coordinate axes the joint probability distribution function (PDF) \( P(\gamma_1, \gamma_2) \) obeys:
\[ P(\gamma_1, \gamma_2) = P(|\gamma_1|, |\gamma_2|) = P(|\gamma_2|, |\gamma_1|). \]

2.2 Convergence: \( \kappa \)

Using eq.(1) and eq.(4) one can show (Bernardeau et al., 1997, and Kaiser, 1998) that the convergence along a given line of sight is:
\[ \kappa \simeq -\frac{\Omega_m}{2} \int_0^{\chi_s} d\chi \ w(\chi, \chi_s) \delta(\chi), \]
with:
\[ w(\chi, \chi_s) = \frac{H_0^2}{c^2} \frac{D(\chi)D(\chi_s - \chi)}{D(\chi_s)} (1 + z), \]
where \( z \) corresponds to the radial distance \( \chi \). Thus the convergence, \( \kappa \), can be expressed very simply as a function of the density field; it is merely an average of the local density contrast along the line of sight. Then, we can see from eq.(8), that there is a minimum value, \( \kappa_{\text{min}}(z_s) \), for the convergence of a source located at redshift \( z_s \), which corresponds to an “empty” beam between the source and the observer (\( \delta = -1 \) everywhere along the line of sight):
\[ \kappa_{\text{min}} = -\frac{3\Omega_m}{2} \int_0^{\chi_s} d\chi w(\chi, \chi_s). \]

Following Valageas (2000a, b) it is convenient to define the “normalized” convergence, \( \hat{\kappa} \), by:
\[ \hat{\kappa} = \frac{\kappa}{\kappa_{\text{min}}} = \int_0^{\chi_s} d\chi \hat{w} \delta, \quad \text{with} \quad \hat{w} = \frac{w(\chi, \chi_s)}{\int_0^{\chi_s} d\chi w(\chi, \chi_s)} \]
which obeys \( \hat{\kappa} \geq -1 \). Here we introduced the “normalized selection function,” \( \hat{w}(\chi, \chi_s) \). One interest of working with normalized quantities like \( \hat{\kappa} \) is that most of the cosmological dependence (on \( \Omega_m, \Omega_L \) and \( z_s \)) and the projection effects are encapsulated within \( \kappa_{\text{min}} \), while the statistics of \( \hat{\kappa} \) (e.g., its PDF) mainly probe the deviations from Gaussianity of the density field which arise from the non-linear dynamics of gravitational clustering. If one smooths the observations with a top-hat window in real space of small angular radius, \( \theta_s \), one rather considers the filtered normalized convergence \( \hat{\kappa}_s \) (where the subscript “s” refers to “smoothed”):
\[ \hat{\kappa}_s = \int_0^{\theta_s} \frac{d\vartheta}{\vartheta} \int_{\vartheta}^{\chi_s} d\chi \hat{w}(\chi, \chi_s) \delta(\chi, D\vartheta). \]
Here \( \vartheta \) is a vector in the plane perpendicular to the line of sight (we restrict ourselves to small angular windows) over which we integrate within the disk \( |\vartheta| \leq \theta_s \); we note this by the short notation \( \int^{\theta_s} \). Thus \( \chi \) is the radial coordinate, while \( D\vartheta \) is the two-dimensional vector of transverse coordinates. Eq.(12) clearly shows that the convergence \( \hat{\kappa}_s \) is actually an average of the density contrast over the cone of angular radius \( \theta_s \).

In the following, it will be convenient to work in Fourier space. Thus, we define the Fourier transform of the density contrast by:
\[ \delta(x) = \int dk e^{ik \cdot x} \delta(k) \]
where \( x \) and \( k \) are comoving coordinates. Then, eq.(12) also reads:
\[ \hat{\kappa}_s = \int_0^{\chi_s} d\chi \hat{w}(\chi, \chi_s) \int dk e^{ik \cdot x} W(k \cdot \vartheta_s) \delta(k), \]
where \( k_\parallel \) is the component of \( k \) parallel to the line of sight and \( k_\perp \) is the two-dimensional vector formed by the components of \( k \) perpendicular to the line of sight. Here we introduced the Fourier form \( W(k \cdot \vartheta_s) \) of the real-space top-hat filter of angular radius \( \theta_s \):
\[ W(k \cdot \vartheta_s) = \int_0^{\theta_s} \frac{d\vartheta}{\vartheta} e^{ik_\parallel \cdot \vartheta} \frac{2J_1(k_\perp \vartheta_s)}{k_\parallel}, \]
where \( J_1 \) is the Bessel function of the first kind of order 1. If we choose another filter (e.g., a Gaussian window rather than a top-hat) the expression (14) remains valid and we simply need to use the relevant Fourier window \( W(k \cdot \vartheta_s) \).

2.3 Shear: \( \gamma \)

For the shear, \( \gamma \), defined in eq.(6), we can obtain expressions which are similar to eqs.(12)-(14). Thus, from eq.(1) and eq.(6) we obtain in Fourier space for the normalized shear \( \hat{\gamma} = \gamma / \kappa_{\text{min}} \):
\[ \hat{\gamma} = \int_0^{\chi_s} d\chi \hat{w}(\chi, \chi_s) \int dk e^{ik \cdot x} \frac{k_\perp^2 - k_\parallel^2 + 2ik_1 k_2}{k_1^2 + k_2^2} \delta(k) \]
which of course are the components of \( k \) along the two orthogonal axes perpendicular to the line of sight which also define the directions (1, 2) used in Subsection 2.1 for the components \( \Phi_{i,j} \) of the shear tensor. In eq.(16) we used again the small-angle approximation. Of course, for the smoothed shear \( \hat{\gamma}_s \) we get, in a similar fashion to eq.(14):
\[ \hat{\gamma}_s = \int_0^{\chi_s} d\chi \hat{w} \int dk e^{ik \cdot x} \frac{k_\perp^2 - k_\parallel^2 + 2ik_1 k_2}{k_1^2 + k_2^2} \]
\times W(k \cdot \vartheta_s) \delta(k). \]

Going back to real-space, we also obtain from eq.(17) for a top-hat filter the expression:
\[ \hat{\gamma}_s = -\int_0^{\chi_s} d\chi \hat{w} \int_{\theta_s}^{\infty} \frac{d\vartheta}{\vartheta} \int_0^{\pi} d\alpha \ e^{i\alpha} \delta(\chi, D\vartheta) \]
where \( \vartheta \) is a vector in the plane perpendicular to the line of sight of length \( \vartheta \) and which makes the polar angle, \( \alpha \), with
the 1-axis (the 2-axis has $\alpha = \pi/2$). Note that we integrate over $\delta$ over all space outside of the disc of radius $\theta_s$. Thus, the comparison with eq. (12) shows that while the smoothed convergence only depends on the matter within the cone formed by the angular window, $\theta_s$, the smoothed shear only depends on the matter outside this cone.

3 CUMULANTS, PDFS AND MODELS FOR THE MANY-BODY DENSITY CORRELATIONS

In this paper, we wish to evaluate the probability distribution function, PDF, of the shear, $\gamma$. To this order, following the approach developed in Valageas (2000a, b) and Munshi & Jain (2000, 2001), we shall first compute the moments (or more precisely the cumulants) of the shear and then obtain the associated PDF. Therefore, we first recall the standard relationship between the PDF and the moment and cumulant generating functions.

3.1 Generating functions

The statistical properties of a random variable $s$ can be obtained from its moment or cumulant generating functions, which are widely used in statistics (see also Ballan & Schaeffer, 1989). Thus, the PDF $P(s)$ can be derived from the generating function, $\psi(y)$, defined from the moments, $\langle s^p \rangle$, (provided they are finite) by:

$$\psi(y) = \sum_{p=0}^{\infty} \frac{(-1)^p}{p!} \langle s^p \rangle y^p \quad (19)$$

since we have the inverse Laplace transform:

$$P(s) = \frac{1}{2\pi i} \int_{-\infty}^{\infty} ds e^{-sy} \psi(y). \quad (20)$$

Indeed, the generating function, $\psi(y)$, is simply the Laplace transform of the PDF $P(s)$:

$$\psi(y) = \int ds e^{-sy} P(s). \quad (21)$$

One can check that expanding the exponential in eq. (21) reproduces eq. (19). If the moments, $\langle s^p \rangle$, diverge one can still define $\psi(y)$ from eq. (21) but the expansion at $y = 0$ of the generating function becomes singular. In practice, one usually introduces the generating function, $\Phi(y)$, defined from the cumulants, $\langle s^p \rangle_c$, by:

$$\Phi(y) = \sum_{p=1}^{\infty} \frac{(-1)^p}{p!} \langle s^p \rangle_c y^p = \ln [\psi(y)]. \quad (22)$$

As is well-known, $\Phi(y)$ is simply the logarithm of $\psi(y)$. For our purposes, we shall actually use the normalized generating function, $\varphi(y)$, defined by:

$$\varphi(y) = \sum_{p=1}^{\infty} \frac{(-1)^p}{p!} \frac{\langle s^p \rangle_c}{\xi^p} y^p \quad \text{with} \quad \xi = \langle s^2 \rangle_c. \quad (23)$$

The comparison of eq. (23) with eq. (22) yields $\psi(y) = \exp[-\varphi(y\xi)/\xi]$, so that eq. (20) now reads:

$$P(s) = \int_{-\infty}^{+\infty} \frac{dy}{2\pi i \xi} e^{[sy - \varphi(y)]/\xi}. \quad (24)$$

As seen in Section 2, the convergence, $\kappa$, or the shear, $\gamma$, are given by a linear integral along the line of sight over the density field. In other words, they can be seen as the superposition of independent layers which each contribute to $\kappa$ or $\gamma$. Indeed, within the small-angle approximation the correlations of the density field will only appear within the transverse directions. Then, the direct calculation of the PDF of such a sum over the redshift of the lenses would yield an infinite number of convolution products which makes it intractable. By contrast, the cumulant generating functions simply add when different layers are superposed (since Laplace transforms change convolutions into ordinary products). This property makes the generating function $\varphi(y)$ a convenient tool to deal with projection effects along the line of sight. This remark is the basis of the method introduced in Valageas (2000a, b). Moreover, the expansion (23) provides a simple way to compute $\varphi(y)$ from the moments (or cumulants) of the random variable one is interested in. This approach can be used for any quantity which is linear over the density field, like the convergence, $\kappa$, studied in Valageas (2000a, b), the aperture-mass, $M_{ap}$, investigated in details in Bernadeau & Valageas (2000) or the shear, $\gamma$, we focus on in this paper.

3.2 PDF for the density field

It is clear that in order to derive the PDF of the shear, $\gamma$, we need the properties of the underlying density field. A first measure of the statistical properties of the density field is given by the PDF $P(\delta_R)$ of the density contrast within spherical cells of radius $R$ and volume $V$:

$$\delta_R = \int_V dx \delta(x). \quad (25)$$

Note that the cumulants, $\langle \delta_R^p \rangle_c$, can also be written in terms of the many-body connected correlation functions, $\xi_p(x_1,..,x_p)$, (as, e.g., Peebles, 1980):

$$p \geq 2 : \langle \delta_R^p \rangle_c = \xi_p = \int_{V^p} dx_1...dx_p \xi_p(x_1,..,x_p), \quad (26)$$

with:

$$\xi_p(x_1,..,x_p) = \langle \delta(x_1) \delta(x_p) \rangle_c. \quad (27)$$

Next, we define the associated normalized generating function, $\varphi(y)$, introduced in eq. (23) and its coefficients, $S_p$, by:

$$\varphi(y) = \sum_{p=2}^{\infty} \frac{(-1)^{p-1}}{p!} S_p y^p, \quad S_p = \frac{\xi_p}{\xi_2}, \quad S_2 = 1, \quad (28)$$

which also yields $P(\delta_R)$ through eq. (24):

$$P(\delta_R) = \int_{-\infty}^{+\infty} \frac{dy}{2\pi i \xi_2} e^{[y\delta - \varphi(y)/\xi_2].} \quad (29)$$

Since the convergence, $\kappa_\gamma$, is the average of the density contrast, $\delta$, with a smooth positive weight over the relevant cone, its cumulants can be estimated from those of the density field in a robust and simple way, by neglecting their geometrical dependence. Therefore, one merely uses the properties of the density field, $\delta_R$, smoothed over spherical cells...
and does not take into account the detailed angular dependence of the many-body correlation functions. As shown in Valageas (2000b), this procedure is indeed sufficient to recover the properties of the convergence. In fact, up to a good accuracy, one may even express the PDF, $P(\kappa_i)$, directly in terms of the PDF, $P(\delta_R)$, of the density contrast, by using the approximation $\phi_{\kappa_i}(y) \simeq \phi(y)$ (see Valages 2000b). Here $\phi$ (also defined in eq.(28)) is the normalized generating function associated with the density contrast $\delta_R$ at the scale and redshift which contribute most to the signal, whereas $\phi_{\kappa_i}$ is the normalized generating function associated with the normalized convergence, $\kappa_i$, as defined by eq.(23). This approximation clearly shows that the “normalized convergence,” $\kappa_i$, mainly probes the non-Gaussianities of the density field (i.e., its very structure) brought by the non-linear gravitational dynamics, while the overall amplitude of the fluctuations (which gives the dependence on the cosmological parameters and the source redshift) is encoded within $\kappa_{\text{min}}$, as we announced earlier, below eq.(11).

3.3 Models for the many-body density correlations

For more intricate filters, this simple approximation is no longer sufficient. This is clearly the case for the shear, $\gamma$, since eq.(18) shows that the relevant volume is now the exterior of the cone of radius $\theta$, and has a geometry which is significantly different from spherical cells. More importantly, the factor $e^{2\alpha_0}$ yields a weight of mean zero and which is both positive and negative and varies with the direction in the transverse plane. This implies that one cannot evaluate the cumulants of $\gamma$ from these moments of the density field over spherical cells; the natural estimate would be zero, since the angular mean of the weight, $e^{2\alpha_0}$, vanishes. The same problem appears for the study of the aperture mass, $M_{\text{ap}}$, which also involves a compensated filter. Therefore, we need an explicit model for the angular dependence of the many-body correlations.

As shown in Bernardeau & Valageas (2000), in the case of the aperture mass, $M_{\text{ap}}$, good results can be obtained at small angles by using a “minimal tree-model” for the connected correlations of the density field. Let us first recall that a general “tree-model” is defined by the property (Schaeffer, 1984, and Groth & Peebles, 1977):

$$\xi_p(\mathbf{x}_1, \ldots, \mathbf{x}_p) = \sum_{(\alpha)} Q^p_\alpha \prod_{t_\alpha} \xi_{t_\alpha}(\mathbf{x}_i, \mathbf{x}_j)$$

where $(\alpha)$ is a particular tree-topology connecting the $p$ points without making any loop, $Q^p_{(\alpha)}$ is a parameter associated with the order of the correlations and the topology involved, $t_{\alpha}$ is a particular labeling of the topology, $(\alpha)$, and the product is made over the $(p-1)$ links between the $p$ points with two-body correlation functions. We show in Fig.1 the three topologies which appear within this framework for the 5-point connected correlation function.

3.3.1 Minimal tree model

A peculiar case of the tree-models described by (30) is the minimal tree-model where the weights $Q^p_{(\alpha)}$ are given by (Bernardeau & Schaeffer, 1992):

$$Q^p_{(\alpha)} = \prod_{\text{vertices of (\alpha)}} \nu_q$$

where $\nu_q$ is a constant weight associated to a vertex of the tree topology with $q$ outgoing lines. Then, for an arbitrary real-space filter, $F(\mathbf{x})$, which defines the random variable $s$ as:

$$s = \int d\mathbf{x} \ F(\mathbf{x}) \ \delta(\mathbf{x}) \ \ \text{and} \ \ \xi_s = \langle s^2 \rangle,$$

it is possible to obtain a simple explicit expression for the normalized generating function, $\phi_s(y)$ (see Bernardeau & Schaeffer, 1992, and Jannink & Des Cloiseaux, 1987):

$$\phi_s(y) = y \int d\mathbf{x} \ F(\mathbf{x}) \ \left[ \xi(T(\mathbf{x})) - \frac{\tau(\mathbf{x}) \xi'(T(\mathbf{x}))}{2} \right]$$

$$\tau(\mathbf{x}) = -\frac{y}{\xi_s} \int d\mathbf{x}' \ F(\mathbf{x}') \ \xi_2(\mathbf{x}, \mathbf{x}') \ \xi'(T(\mathbf{x}'))$$

where the function $\zeta(\tau)$ is defined as the generating function for the coefficients $\nu_p$:

$$\zeta(\tau) = \sum_{p=1}^{\infty} \frac{(-1)^p}{p!} \nu_p \ \tau^p \ \text{with} \ \nu_1 = 1.$$
because the 2-d compensated filter $F(\vec{b})$ now strongly depends on the polar distance $|\vec{b}|$ (in the 2-d transverse plane). Such a study is described in detail in Bernardeau & Valageas (2000) where it is shown that this minimal tree-model provides good results when compared with numerical simulations.

Of course, we can apply the same procedure, developed in Bernardeau & Valageas (2000) for $M_{ap}$, to the shear, $\gamma$. We simply need to use the corresponding real-space filter, $P$. From eq.(18) we see that this is $F(\vec{b}) = -H(\vec{b} > \theta_s) \cos(2\alpha)/\pi b^2$ for $\gamma_1$ and $F(\vec{b}) = -H(\vec{b} > \theta_s) \sin(2\alpha)/\pi b^2$ for $\gamma_2$, where $H$ is the Heaviside function with obvious notations. Then, the resummation (33)-(34) yields a two-dimensional non-linear implicit equation which is not very convenient for numerical purposes. Therefore, it is interesting in this case to investigate whether a simpler model for the many-body correlations of the density field could still provide good predictions for the PDF, $P(\gamma_1)$, of the shear component, $\gamma_1$.

### 3.3.2 Stellar model

Thus, we introduce in this paper a “stellar-model”, which is a particular case of the tree-models defined in (30), where we only keep the stellar diagrams (e.g., the graph (a) in Fig.1 for the 5-point connected correlation). Thus, the point-point connected correlation $G_\rho$ of the density field can now be written as:

$$
\xi_\rho(x_1, \ldots, x_p) = \sum_{i=1}^p \prod_{j \neq i} \xi_2(x_i, x_j).
$$

(38)

The advantage of the stellar-model (38) is that it leads to very simple calculations in Fourier space. Thus, let us define the power-spectrum, $P(k)$, of the density contrast by:

$$
\langle \delta(k_1) \delta(k_2) \rangle = \delta_D(k_1 + k_2) P(k_1),
$$

(39)

where $\delta_D$ is Dirac’s distribution, whence:

$$
\xi_2(x) = \int dk e^{i k \cdot x} P(k).
$$

(40)

Then, eq.(38) reads in Fourier space:

$$
\langle \delta(k_1) \delta(k_2) \rangle = \sum_{i=1}^p \prod_{j \neq i} P(k_j).
$$

(41)

Of course, the Dirac factor, $\delta_D(k_1 + \ldots + k_p)$, simply translates the fact that the many-body correlations are invariant through translations. Therefore, in the following we shall use eq.(41) for the connected correlations of the density field which appear in the calculation of the cumulants of the shear, $\gamma$.

### 3.4 Amplitude of the density correlations

The stellar-model introduced in Section 3.3.2 provides the angular dependence of the many-body correlations of the density field. However, we still need to determine their overall amplitude, that is the non-linear power-spectrum, $P(k)$, and the coefficients, $S_p$, which enter eqs.(38)-(41). First, we obtain the non-linear evolution of the power-spectrum from the prescription given by Peacock & Dodds (1996), which is also a good fit to numerical simulations. Second, we determine the coefficients, $S_p$, as follows. We characterize the amplitude of the density correlations at scale $R$ and redshift $z$ through the parameters $S_p$, defined in eq.(28), which give the amplitude of the many-body correlations over spherical cells of radius $R$. As seen from eq.(28)-(29) they also define the PDF, $P(\delta_R)$. We could obtain these coefficients, $S_p$, through a parameterization of their generating function, $\varphi(y)$. However, we prefer to specify $\varphi(y)$ through a function $\zeta(\tau)$ which is defined by the implicit system (36)-(37). Note that this procedure is merely a convenient parameterization of $\varphi(y)$ and it does not assume by itself a minimal tree-model. The results recalled in Section 3.3.1 simply mean that in case the actual many-body correlations exactly obey a minimal tree-model, the function $\zeta(\tau)$ from eqs.(36)-(37) could also be interpreted as a good approximation to the (slightly different) function $\zeta_\nu(\tau)$ defined in eq.(35) as the generating function of the vertices $v_p$.

The reason we prefer to parameterize the coefficients, $S_p$, through $\zeta(\tau)$ rather than directly through $\varphi(y)$ is that this procedure was already checked to provide good results in previous works. Thus, it was seen in Bernardeau & Schneider (1992) to give good predictions for the galaxy and matter correlations, while in Bernardeau & Valageas (2000) it was shown to yield good predictions for the PDF, $P(M_{ap})$, of the aperture-mass. Since the aperture-mass is closely related to the shear, $\gamma$, we can expect this procedure to give good results for the PDF, $P(\gamma)$, too. Note however that the calculations developed in Bernardeau & Valageas (2000) used the minimal tree-model recalled in Section 3.3.1 rather than the stellar-model of Section 3.3.2 for the angular dependence of the many-body correlations. Nevertheless, we checked numerically that both models actually give close results for $P(M_{ap})$. Thus, as in Bernardeau & Valageas (2000) we parameterize $\zeta(\tau)$ as:

$$
\zeta(\tau) = \left(1 + \frac{\tau}{\kappa} \right)^{-\kappa} - 1,
$$

(42)

where we have kept the usual notation, $\kappa$ (not to be confused with the convergence), for the free parameter which enters the definition of $\zeta(\tau)$.

Next, we choose the value of this parameter, $\kappa$, so as to reproduce the skewness, $S_3$, in both the highly non-linear and quasi-linear regimes, using the relation:

$$
\kappa = \frac{3}{S_3 - 3}.
$$

(43)

For $S_3 > 3$ we have $\kappa > 0$ while for $0 < S_3 < 3$ we have $\kappa < -1$. In the limit $S_3 \to 3$ the function $\zeta(\tau)$ defined in eq.(42) goes to $\zeta(\tau) \to e^{-\tau} - 1$ (and $|\kappa| \to \infty$). We take for the skewness in the highly non-linear regime, $S_3^{NL}$, the prediction of HEPT (Scoccimarro & Frieman, 1999) and for the quasi-linear regime, $S_3^{QL}$, the exact result obtained from perturbative theory:

$$
S_3^{NL} = 3 \frac{4 - n^2}{1 + 2n + n^2}, \quad S_3^{QL} = 34 \frac{n}{7} - (n + 3).
$$

(44)

Here $n$ is the local slope of the linear power-spectrum at the typical wavenumber $k_0$ probed by the observations at redshift $z$. From eq.(14) and eq.(17), we define this typical wavenumber as:

$$
k_0(z) = \frac{1}{D(z)\theta_s}.
$$

(45)
Finally, we introduce the power, \( \Delta^2(k) \), per logarithmic interval defined as:
\[
\Delta^2(k, z) = 4\pi k^3 P(k, z).
\]
(46)
Then, for intermediate regimes defined by \( 1 < \Delta^2(k, z) < \Delta_{vir}(z) \), we use the simple linear interpolation:
\[
S_3(z) = S_{3L} + \frac{\Delta^2(k, z) - 1}{\Delta_{vir}(z) - 1} (S_{3L} - S_{3QL}).
\]
(47)
For \( \Delta^2(k, z) < 1 \) (i.e. the quasi-linear regime), we take \( S_3 = S_{3QL} \), while for \( \Delta^2(k, z) > \Delta_{vir}(z) \) (i.e., the highly non-linear regime), we take \( S_3 = S_{3NL} \). Here \( \Delta_{vir} \) is the density contrast at virialization given by the usual spherical collapse (thus \( \Delta_{vir} \sim 178 \) for a critical-density universe). Indeed, the threshold \( \Delta^2 > \Delta_{vir} \) describes the highly non-linear regime where most of the matter at scale \( 1/k \) has collapsed into non-linear structures.

Finally, as can be seen from eq.(41), the coefficients \( S_p \) and \( \tilde{S}_p \) are related by:
\[
S_p = \tilde{S}_p \int_0^1 dt \, 3t^2 \left[ \frac{\int B R}{\int B^2} P(kR) \right]^{p-1}.
\]
where we introduced the Fourier transform, \( F(kR) \), of a 3-d top-hat of radius \( R \):
\[
F(kR) = \int \frac{dx}{V^3} e^{ikx} = \frac{3 \sin(kR) - (kR) \cos(kR)}{(kR)^3}.
\]
However, in the following we shall use the simple approximation:
\[
\tilde{S}_p \simeq S_p.
\]
(50)
Alternatively, we may define the function \( \varphi(y) \) obtained from (36)-(37) and our choice of \( \zeta(\tau) \) as the generating function of the coefficients \( \tilde{S}_p \), rather than \( S_p \), through its Taylor expansion at \( y = 0 \).

It is interesting to note that the implicit system (36)-(37) appears naturally in the quasi-linear regime. Indeed, as shown in Bernardeau (1994) and Valageas (2002a), this regime the generating function \( \varphi(y) \) exactly obeys the relation (36)-(37), where \( \zeta(\tau) \) is closely related to the spherical dynamics for the non-linear density contrast. However, the latter is of the form (42) only in the limit \( \Omega_m \to 0 \) with no smoothing (then, the parameter \( \kappa \) is equal to 3/2). Nevertheless, the form (42) still provides a good approximation for more general cases where \( \kappa \) is adjusted so as to recover the exact skewness. Here we note that the system (36)-(37) usually yields a branch cut along the negative real axis for \( y < y_* \), with \( y_* \sim -0.1 \) for \( \varphi(y) \); see Bernardeau & Schaeffer (1992).

For the function \( \zeta(\tau) \) given in eq.(42) one gets (for \( \kappa > 0 \) or \( \kappa < -2 \), that is \( S_3 > 3/2 \)):
\[
y_* = -\frac{\kappa}{k + 2} \left( \frac{\kappa + 1}{\kappa + 2} \right)^{k+1} \quad \text{and} \quad \tau_* = -\frac{\kappa}{k + 2}.
\]
(51)
Such a singularity leads to an exponential cutoff for the PDF, \( \mathcal{P}(\delta_R) \sim e^{y_* \delta_R^2/2} \) at large densities. As shown in Valageas (2002a), in the case of the quasi-linear regime this is actually an artefact and one must go through \( y_* \) up to a second branch for \( \varphi(y) \), which yields a large density cutoff for \( \mathcal{P}(\delta_R) \) which is shallower than a simple exponential.

However, it was seen by numerical computations in Valageas (2002a) that this feature does not change the shape of the PDF by much in the range of interest. Moreover, in this paper we are mostly interested in the non-linear regime (which largely dominates for \( \theta < 4^\circ \) and \( z_s = 1 \)) where the singularity \( y_* \) is meaningful and the PDF shows indeed an exponential cutoff (since it is actually defined from the coefficients \( S_p \)).

Finally, we must point out that the simple model developed above has no new free parameter. Indeed, the power-spectrum evolution is obtained from Peacock & Dodds (1996), while all many-body correlations are defined by the stellar-model (38) with their amplitude given by the sole coefficient \( S_3 \) (i.e. the skewness of \( \mathcal{P}(\delta_R) \)) which is exact in the quasi-linear regime and obtained from HEPT (Scoccimarro & Frieman, 1999) in the non-linear regime. In the form given here, our model is obviously consistent with the “stable-clustering Ansatz” (e.g., Peebles 1980). However, it could incorporate possible deviations from the stable-clustering Ansatz through the non-linear power-spectrum, \( P(k) \), or through an additional dependence on redshift and scale for the skewness, \( S_3 \), in the non-linear regime.

4 COMPUTATION OF THE PDFS

4.1 PDF of the shear components, \( \gamma_1 \) and \( \gamma_2 \)

We can now use the model developed in Section 3 to describe the density field in order to compute the PDF, \( \mathcal{P}(\gamma_1) \), of the shear component, \( \gamma_1 \). First, we obtain from eq.(17):
Then, the cumulant of order $c$ scale, shift $z$ is
\[
\delta(k_{\perp 1} + \ldots + k_{\perp p}) = \int \frac{dt}{(2\pi)^2} e^{it(k_{\perp 1} + \ldots + k_{\perp p})},
\]
so that eq.(55) reads:
\[
\langle \gamma_{1c}^P \rangle_c = \int \frac{d\chi}{2\pi} \tilde{S}_p (2\pi \hat{w})^p \int \frac{dt}{(2\pi)^2} \int dk_1 \ldots dk_p
\]
\[
\times \prod_{j=1}^p \cos 2\phi_j W(k_j D\theta_s) e^{i t k_j (\tilde{w} \cos \phi_j - \phi)} \times P(k_2) \ldots P(k_p),
\]
where we have dropped the subscript, $\perp$, for the 2-d transverse wavenumbers $k_{\perp}$, and we have noted $\phi$ the polar angle of the vector $t$. The integration over the angles $\phi_j$ yields:
\[
\int_0^{2\pi} d\phi_j \cos 2\phi_j e^{i t k_j (\tilde{w} \cos \phi_j - \phi)} = -2\pi \cos 2\phi J_2(tk_j)
\]
so that eq.(57) becomes:
\[
\langle \gamma_{1c}^P \rangle_c = \int \frac{d\chi}{2\pi} \tilde{S}_p (2\pi \hat{w})^p \int \frac{dt}{(2\pi)^2} \int \frac{d\phi_j}{2\pi} (-\cos 2\phi)^p
\]
\[
\times \prod_{j=1}^p \int_0^{2\pi} d\phi_j \tilde{S}_k W(k_j D\theta_s) J_2(tk_j) I_{\gamma_1}(t, \chi) \delta(\phi_j - \phi) I_{\gamma_1}(t, \chi)^p - 1
\]
where we have made the change of variable $t \rightarrow t D\theta$ and we have introduced $I_{\gamma_1}(t, \chi)$ defined by:
\[
I_{\gamma_1}(t, \chi) = 4\pi^2 \int_0^{2\pi} dk W(k D\theta_s) J_2(tk D\theta_s) P(k)
\]
\[
= \pi \int_0^{2\pi} \frac{dt}{k} \frac{\Delta^2(k, z)}{k} W(k D\theta_s) J_2(tk D\theta_s).
\]
For the angular top-hat filter, $W$, given by eq.(15) in Fourier space, we can perform the integration over $k_1$ in eq.(59), using the standard properties of Bessel functions, and we obtain:
\[
\langle \gamma_{1c}^P \rangle_c = \int \frac{d\chi}{2\pi} \tilde{S}_p \hat{w}^p \int \frac{dt}{t} \int_0^{2\pi} \frac{d\phi}{\gamma_{1c}} \tilde{S}_p (\hat{w} \cos 2\phi)^p I_{\gamma_1}^{p-1}.
\]
Here we have dropped the minus sign in front of the factor $\cos 2\phi$ because all odd moments vanish, as seen by integration over $\phi$. Thus, we recover the fact that the PDF, $P(\gamma_1)$, is even, which was obvious from the definition of the shear since there is no preferred direction in the system.

Finally, the expression (61) allows us to obtain the generating function $\varphi_{\gamma_1}(y)$ of the shear component $\gamma_{1s}$, defined as in eq.(24), using the resummation (28) (where we replace $S_p$ by $\tilde{S}_p$; see eq.(50)):
\[
\varphi_{\gamma_1}(y) = \int_0^{\chi_2} d\chi \int_0^{\chi_1} \frac{dt}{t} \int_0^{2\pi} \frac{d\phi}{\gamma_{1s}} \tilde{S}_p (\hat{w} \cos 2\phi)^p I_{\gamma_1}^{p-1}
\]
\[
\times \varphi (y \cos 2\phi I_{\gamma_1} \hat{w} \cos 2\phi) \xi_{\gamma_1}. \quad \gamma_{1s}
\]
Here we have introduced the variance $\xi_{\gamma_1} = \langle \gamma_{1s}^2 \rangle$. From eq.(61) it is given by:
\[
\xi_{\gamma_1} = \int_0^{\chi_2} d\chi \int_0^{\chi_1} \frac{dt}{t} \int_0^{2\pi} \frac{d\phi}{\gamma_{1s}} (\hat{w} \cos 2\phi)^2 I_{\gamma_1}
\]
\[
= \frac{1}{2} \int_0^{\chi_2} d\chi \hat{w} \pi \int_0^{2\pi} \frac{dt}{k} \frac{\Delta^2(k, z)}{k} W(k D\theta_s)^2.
\]
Here we have used $\tilde{S}_2 = 1$; see eq.(41), which is still verified by our parameterization (42). Of course, eq.(63) is exact,
within the small-angle approximation, and the last equality in eq.(63) shows that we recover the property \( \tilde{\gamma}_s = \hat{\kappa}_s^2 \).

For the shear component, \( \gamma_2 \), we obtain the same equations with \( \cos 2\theta \) replaced by \( \sin 2\theta \). This leaves the cumulants \( \langle \gamma_2^0 \rangle_c = \langle \gamma_2^2 \rangle_c \) unchanged and we get the same generating function. Therefore, we recover \( P(\gamma_1) = P(\gamma_2) \). Finally, the PDF is given by the inverse Laplace transform (24):

\[
P(\tilde{\gamma}_{1s}) = \int_{-i\infty}^{+i\infty} \frac{dy}{2\pi i \xi_{1s}} e^{[\tilde{\gamma}_{1s}y - \varphi_{1s}(y)]/\xi_{1s}}.
\]

(64)

The generating function \( \varphi(y) \) obtained in Section 3.4 shows a branch cut along the negative real axis for \( y < y_s \); see eq.(51). Then, we see from eq.(62) that the generating function \( \varphi_{1s}(y) \) for the shear component \( \gamma_1 \) shows two symmetric branch cuts along the real axis, for \( y < -y_s \gamma_1 \) and \( y > y_s \gamma_1 \), with:

\[
y_{s, \gamma_1} = \min_{z,t} |y_s \xi_{1s} I_1 w|.
\]

(65)

Let us recall here that the generating function \( \varphi(y;z) \) explicitly depends on the redshift \( z \) along the line of sight through the skewness \( S_3(z) \) of the density contrast at the typical wave number \( k_s(z) \) probed at this redshift, see Section 3.4. Hence its singularity \( y_s(z) \) also depends on the redshift \( z \). Next, the singularities of \( \varphi_{1s}(y) \) at \( \pm y_s \gamma_1 \) also imply an exponential tail for \( P(\tilde{\gamma}_{1s}) \) at large \( \tilde{\gamma}_{1s} \):

\[
|\tilde{\gamma}_{1s}| \gg 1 : \quad P(\tilde{\gamma}_{1s}) \sim e^{-y_s \gamma_1 |\tilde{\gamma}_{1s}|/\xi_{1s}}.
\]

(66)

which translates into similar exponential tails for the shear \( \gamma_{1\theta} = |\kappa_{\min}| \gamma_{1s} \).

4.2 PDF of the shear modulus, \( |\gamma| \)

From the smoothed shear components, \( \tilde{\gamma}_{1s} \) and \( \tilde{\gamma}_{2s} \), we may define the smoothed modulus, \( |\tilde{\gamma}| \), as:

\[
|\tilde{\gamma}| = |\tilde{\gamma}_{1s} + i\tilde{\gamma}_{2s}| = \sqrt{\tilde{\gamma}_{1s}^2 + \tilde{\gamma}_{2s}^2}.
\]

(67)

Note that eq.(67) means that we take the modulus of the shear after smoothing of the two components, \( \gamma_1 \) and \( \gamma_2 \). Since smoothing and taking the modulus do not commute, the quantity \( |\tilde{\gamma}| \) defined in eq.(67) is not the smoothed shear obtained by first taking the modulus \( |\gamma| = \gamma_1 + i\gamma_2 \) at each point on the sky and second applying a top-hat smoothing to this map. In this paper we shall only consider the smoothed modulus \( |\tilde{\gamma}| \) as defined in eq.(67). Indeed, since taking the modulus is a non-linear process one cannot directly apply the method used in this paper to derive the PDF of the modulus \( \gamma \) defined by first taking the modulus of the shear. By contrast, we can easily study the smoothed modulus \( |\tilde{\gamma}| \) defined in eq.(67) because it can be directly obtained from the properties of the smoothed shear components, \( \tilde{\gamma}_{1s} \) and \( \tilde{\gamma}_{2s} \). Indeed, from eq.(67) we obtain the relation:

\[
P(|\tilde{\gamma}|) \equiv \int_0^{2\pi} d\alpha \, P(|\tilde{\gamma}| \cos \alpha, |\tilde{\gamma}| \sin \alpha),
\]

(68)

where we have introduced the joint PDF, \( P(\tilde{\gamma}_{1s}, \tilde{\gamma}_{2s}) \), of the two shear components, \( \tilde{\gamma}_{1s} \) and \( \tilde{\gamma}_{2s} \). Of course, following the method used in Section 4.1 for the shear components we can obtain the joint PDF, \( P(\tilde{\gamma}_{1s}, \tilde{\gamma}_{2s}) \), from the cumulants, \( \langle \tilde{\gamma}_{1s}^2 \tilde{\gamma}_{2s}^2 \rangle_c \). Thus, proceeding along the lines of Section 4.1 the analog of eq.(61) now reads:

\[
\langle \tilde{\gamma}_{1s}^p \tilde{\gamma}_{2s}^p \rangle_c = \int_0^{\chi_s} d\chi \, \tilde{S}_p + p_1 + p_2 \int_1^\infty \frac{dt}{t} \int_0^{2\pi} \frac{d\phi}{\pi} \chi \left( -\cos 2\phi \right)^{p_1} \left( -\sin 2\phi \right)^{p_2} I_{1+1+2-1}.
\]

(69)

Then, the generating function, \( \Phi_{\tilde{\gamma}_{1s}, \tilde{\gamma}_{2s}}(y_1, y_2) \), defined as in eq.(22) by:

\[
\Phi_{\tilde{\gamma}_{1s}, \tilde{\gamma}_{2s}}(y_1, y_2) = \sum_{p_1, p_2=0}^\infty \frac{(-1)^{p_1 + p_2}}{p_1! \, p_2!} \langle \tilde{\gamma}_{1s}^p \tilde{\gamma}_{2s}^p \rangle_c y_1^{p_1} y_2^{p_2}
\]

(70)

can be written from eq.(69) as:

\[
\Phi_{\tilde{\gamma}_{1s}, \tilde{\gamma}_{2s}}(y_1, y_2) = \int_0^{\chi_s} d\chi \int_1^\infty \frac{dt}{t} \int_0^{2\pi} \frac{d\phi}{\pi} \frac{1}{\xi_{1s}} \chi \left( y_1 \cos 2\phi + y_2 \sin 2\phi \right)^p.
\]

(71)

Using the resummation (28) we get:

\[
\Phi_{\tilde{\gamma}_{1s}, \tilde{\gamma}_{2s}}(y_1, y_2) = -\int_0^{\chi_s} d\chi \int_1^\infty \frac{dt}{t} \int_0^{2\pi} \frac{d\phi}{\pi} \frac{1}{\xi_{1s}} \chi \left( y_1 \cos 2\phi + y_2 \sin 2\phi \right)^p \left( y_1 \cos 2\phi + y_2 \sin 2\phi \right) \langle I_{\gamma_1} \rangle_{\xi_{1s}} (\chi, \chi, z).
\]

(72)

Next, from the cumulant generating function, \( \Phi_{\tilde{\gamma}_{1s}, \tilde{\gamma}_{2s}} \), we define the normalized generating function, \( \varphi_{\tilde{\gamma}_{1s}, \tilde{\gamma}_{2s}} \), by:
\[ \varphi_{\gamma_1, \gamma_2}(y_1, y_2) = -\xi_{\gamma_1} \Phi_{\gamma_1, \gamma_2} \left( \frac{y_1}{\xi_{\gamma_1}}, \frac{y_2}{\xi_{\gamma_1}} \right) \]  

(73)

and eq. (72) yields:

\[ \varphi_{\gamma_1, \gamma_2}(y_1, y_2) = \int_0^x d\chi \int_1^{\infty} \frac{dt}{t} \int_0^{2\pi} \frac{d\phi}{\pi} \frac{\xi_{\gamma_1}}{\xi_{\gamma_1}} \times \varphi \left( y_1 \cos 2\phi + y_2 \sin 2\phi \right) \frac{I_{\gamma_1}}{\xi_{\gamma_1}}. \]  

(74)

Then, the joint PDF, \( P(\gamma_{1s}, \gamma_{2s}) \), is given by:

\[ P(\gamma_{1s}, \gamma_{2s}) = \int_{-\infty}^{\infty} d\gamma_1 d\gamma_2 \times [\varphi(\gamma_{1s} + \gamma_{2s} \gamma_{1s} \varphi(\gamma_{1s}, \gamma_{2s}(y_1, y_2))] / \xi_{\gamma_1}. \]  

(75)

It is interesting to compare our results (74) and (75) with the prediction we would obtain if we had assumed that \( \hat{\gamma}_1 \) and \( \hat{\gamma}_2 \) were uncorrelated, but still described by the PDF derived in Section 4.1. Then, the normalized generating function, \( \varphi_{\gamma_1, \gamma_2}(y_1, y_2) \) would simply be given by:

\[ \varphi^\text{uncor}(\gamma_{1s}, \gamma_{2s}) = \varphi(\gamma_{1s}) + \varphi(\gamma_{2s}). \]  

(76)

Substituting this into eq. (75), we obviously recover the property:

\[ P^\text{uncor}(\gamma_{1s}, \gamma_{2s}) = P(\gamma_{1s}) P(\gamma_{2s}). \]  

(77)

Therefore, the fact that \( \varphi_{\gamma_1, \gamma_2} \) obtained in eq. (74) cannot be written as the sum (76) shows that within our stellar model of Section 3.3.2 for the density field, both shear components, \( \gamma_{1s} \) and \( \gamma_{2s} \), are correlated, which is actually quite natural. Moreover, the comparison of (75) with (77) will allow us to study numerically the influence of these correlations onto the PDF of the smoothed modulus \( |\hat{\gamma}_s| \).

Finally, we can note that for a Gaussian density field we simply have \( \varphi(y) = -y^2/2 \), see eqs. (26)-(28). This also corresponds to \( S_2 = 1 \) and \( S_p = 0 \) for \( p \geq 3 \). Hence all our results are exact for the case of a Gaussian density field, where we simply need to use for the generating function: \( \varphi_{\text{Gauss}}(y) = -y^2/2 \). Then, we recover from eq. (62) that the shear components \( \gamma_1 \) and \( \gamma_2 \) are also Gaussian random variables, since we find \( \varphi_{\text{Gauss}}(y) \propto y^2 \). This was actually obvious from the definition of the shear as a linear integral over the density field, see eq. (16). On the other hand, we can see from eq. (74) that in the Gaussian case we have \( \varphi_{\gamma_{1s}, \gamma_{2s}}(y_1, y_2) = \varphi_{\gamma_{1s}}(y_1) + \varphi_{\gamma_{2s}}(y_2) \), since the cross-correlation term \( y_1 \cos 2\phi x y_2 \sin 2\phi \) vanishes after integration over \( \phi \). Thus, we recover the fact that for a Gaussian density field the two shear components \( \gamma_1 \) and \( \gamma_2 \) along the directions (1,2) are uncorrelated. This is not surprising since for a Gaussian density field the Fourier modes \( \delta(k) \) are uncorrelated. This yields a Rayleigh distribution for the shear modulus \( |\gamma_s| \).

5 Lensing statistics from numerical simulations

We compare in the following Section our analytical results for the shear with weak lensing statistics obtained in cosmological N-body simulations. Couchman, Barber & Thomas (1999) developed an algorithm for computing the shear in three dimensions at locations within the simulation volumes and we have applied this algorithm to the simulations of the Hydra Consortium* produced using the 'Hydra' N-body hydrodynamics code (Couchman, Thomas & Pearce, 1995).

We used two simulations, a flat cosmology with a cosmological constant and an open cosmology with zero cosmological constant. These will be referred to as the LCDM and OCDM cosmologies, respectively. Both contained dark matter particles only of mass \( 1.29 \times 10^{11} h^{-1} \) solar masses, where \( h \) expresses the value of the Hubble parameter in units of \( 100 \text{ km s}^{-1} \text{ Mpc}^{-1} \). The number of particles in each cosmology was \( 86^3 \). We used a variable particle softening in the code with a minimum value in box units of 0.0007(1 + z), where \( z \) is the redshift of the simulation volume. The simulation volumes had comoving side-dimensions of \( 100 h^{-1} \text{ Mpc} \).

To avoid obvious structure correlations between adjacent boxes, each was arbitrarily translated, rotated (by multiples of 90°) and reflected about each coordinate axis, and in addition, each complete run was performed 10 times in each cosmology. The cosmological and simulation parameters of both the LCDM and OCDM simulations are given in Table 1.

| \( \Gamma \) | \( \Omega_m \) | \( \Omega_{\Lambda} \) | \( \sigma_8 \) | \( \theta_{\text{res}} \) | \( \theta_{\text{survey}} \) |
|---|---|---|---|---|---|
| LCDM | 0.25 | 0.3 | 0.7 | 1.22 | 0°.34 | 2°.6 |
| OCDM | 0.25 | 0.3 | 0.0 | 1.06 | 0°.37 | 2°.8 |

* (http://hydra.mcmaster.ca/hydra/index.html)
for the lines-of-sight and the locations within the simulations for the computations of the 3-d shear is as described by Barber (2002). In the present work, a total of 455 × 455 lines of sight were used which completely filled the simulation box immediately preceding the one at redshift 1 and allowed regular sampling of the field of view. In addition, 300 regularly-spaced evaluation locations for the shear along each line of sight were specified in each simulation volume to adequately sample the varying gravitational potential. In the LCDM cosmology, the full field of view was 2°.6 × 2°.6 and in the OCDM cosmology, 2°.8 × 2°.8. The angular resolution in the LCDM cosmology was 0°.34, which equates to the minimum value of the particle softening at the optimum redshift, \( z = 0.36 \), for lensing of sources at a redshift of 1. In the case of the OCDM cosmology, the angular resolution was 0°.37. In both cosmologies, to allow for the larger angular size of the minimum softening at low redshifts and also for the range of particle softening scales above the minimum value, we have not declared our results below angular scales of 1 arcminute.

To obtain the lensing statistics for sources at redshift \( z_s = 1 \), we computed the required values assuming the sources were located at the front face of the simulation volume whose redshift was closest to 1. Therefore in the LCDM cosmology, the sources were at \( z_s = 0.99 \) and in the OCDM cosmology they were at \( z_s = 1.03 \).

Once the 3-d shear had been computed at each location, the multiple lens-plane theory, described clearly in Schneider, Ehlers & Falco (1992), was applied along the lines of sight. The procedure involved combining the intermediate Jacobian matrices computed using the appropriate values of the angular diameter distances in the relevant cosmology. Finally, the Jacobian matrices appropriate at \( z = 0 \) for sources at \( z_s = 1 \) for every line of sight were obtained.

From the Jacobian matrices in each of the 10 simulations runs for each cosmology, we computed the shear components, \( \gamma_1 \) and \( \gamma_2 \), and the full shear, \( \gamma \). The values for the smoothed components were computed on angular scales of 1°.0, 2°.0, 4°.0 and 8°.0, using a top-hat filter.

The computed values for the probability distributions from each of the \( N = 10 \) runs in each cosmology were averaged, and the errors on the means of 1σ/√\( N \) determined.

6 COMPARISON OF ANALYTICAL PREDICTIONS WITH RESULTS FROM NUMERICAL SIMULATIONS

We compare here our analytical predictions for the smoothed shear components and for the shear modulus with the results of the numerical simulations. The quantities plotted are for various smoothing angles from 1° to 8° in the two different cosmologies, LCDM and OCDM, and in each case the source redshift is fixed at unity.

Firstly we compare the results for the shear components, plotted in Figure 2, for the LCDM cosmology, and Figure 3, for the OCDM cosmology. Since the PDFs for the components are even functions we only plot the PDF, \( \mathcal{P}(\gamma_1) \), over \( \gamma_1 ≥ 0 \). The solid lines correspond to our analytical predictions, (62) and (64), based on the stellar-model. For comparison, we also plot the Gaussian distribution which has the same variance (the dashed curves), so as to clearly show the deviations from Gaussianity (this corresponds to a Gaussian underlying density field). We find that at small angular scales, \( \theta_s ≤ 4' \), the results from the numerical simulations are very accurately described by our non-linear predictions where there are clear signs of departure from Gaussianity, especially in the tails of the PDF. Of course, this merely reflects the non-linearity of the underlying mass distribution probed by these small angular windows. At larger smoothing angles, \( \theta_s > 4' \), the stellar Ansatz seems to over-predict the non-linearity induced by gravitational clustering and the results lie between the fully non-linear predictions and the Gaussian approximations. This behaviour may be due to the simple interpolation of eq.(47) for the transition between the linear and highly non-linear regimes, which may not be sufficiently accurate. However, we did not try in this work to improve over the simple interpolation (47) because we expect the stellar model to be less accurate in the weakly non-linear regime. Indeed, in the quasi-linear regime the angular dependence of the many-body correlations can be derived from perturbation theory and one can see that for the three-point function it already violates eqs.(38)-(41). On the other hand, in the highly non-linear regime, the angular dependence of the many-body correlations gets shallower and eqs.(38)-(41) actually agree with numerical simulations for the three-point function (Scoccimarro & Frieman 1999).

Of course, at very large scales (which only probe the linear regime) our results would become exact (within the small angle approximation) as we would recover the simple Gaussian.

The PDFs for the shear modulus are presented in Figure 4, for the LCDM cosmology, and Figure 5, for the OCDM cosmology. The solid lines show our analytical predictions (68)-(74)-(75) from the stellar-model. For comparison, we also plot (dotted-dashed curves) the results obtained by assuming that the two shear components are uncorrelated but their own PDFs are still given by our non-linear prediction; this corresponds to eqs.(76)-(77). Therefore, we see from Figures 4 and 5 that the correlation between both components, \( \gamma_1 \) and \( \gamma_2 \), only has a rather small influence on the PDF.
of the modulus, $\mathcal{P}(\gamma_i)$). For reference we have also shown the results for the Gaussian case (dashed curves), where the PDF for $|\gamma|$ is given by a Rayleigh distribution with a variance equal to that of the shear components (dashed curves). Comparison of our predictions with the results of the numerical simulations shows that there is a remarkable agreement in both cosmologies (as found for the shear components), especially at the smaller smoothing scales. At the larger scales, the agreement is less good for the same reasons as described above. The sample variance arising from the various realisations also makes the comparison more difficult; as the size of the smoothing radius is increased, the number of patches with completely independent information decreases, thereby increasing the sample variance. In addition, although we made use of 10 different simulation runs (as described in Section 5), they were all based upon the same N-body realisation in each cosmology, so the separate runs may not contain completely independent information. Much larger simulations would be required to quantify any departure from the analytical predictions.

Finally, we display in Figure 6 our results over a linear-scale for $\mathcal{P}$ rather than the logarithmic scale of the previous Figures, for the LCDM cosmology and the smoothing angle $\theta_s = 1'$. While the previous Figures emphasized the tails of the PDFs we can clearly see in Figure 6 the shapes of the PDFs near their maxima. In agreement with the discussion above, we can check that our predictions match the deviations from Gaussianity observed in the numerical simulations. Thus, the PDF, $\mathcal{P}(\gamma_1)$, of the shear component, $\gamma_1$, is more sharply peaked around $\gamma_1 = 0$ than a Gaussian. Moreover, the location of the maximum of the PDF, $\mathcal{P}(\gamma)$, of the shear modulus is accurately reproduced by our analytical model. These features might be useful, in addition to the extended tails seen in the previous Figures, to measure in weak lensing surveys the departures from Gaussianity brought by the non-linear gravitational dynamics. This may also help to break the degeneracies in the measurement of the cosmological parameters.

7 DISCUSSION

Extending earlier studies we have developed analytical techniques based on a simple hierarchical Ansatz, which we called the stellar model, to compute various statistical quantities associated with cosmic shear. Thus, we have computed the full probability distribution functions associated with the smoothed shear components, $\mathcal{P}(\gamma_i(\theta_s))$, and the modulus of cosmic shear, $\mathcal{P}(|\gamma(\theta_s)|)$. Here $\gamma(\theta_s)$ is defined by first smoothing both shear components and second taking their modulus. We have also derived the joint-PDF, $\mathcal{P}(\gamma_1(\theta_s), \gamma_2(\theta_s))$. These results extend similar studies applied to the convergence field. However, being a spin-2 field, statistics from the shear are much richer than those from the convergence field and, more importantly, can be directly used for observational purposes where the non-trivial survey geometry can pose a considerable problem for reconstructing convergence $\kappa(\theta_s)$ maps.

The comparison of our analytical calculations with the results of the numerical simulations shows that our simple method yields accurate predictions at small angular scales, $\theta_s \leq 4'$ (for a source redshift $z_s = 1$) which probe the non-linear regime for the underlying density field. Therefore, we can use with confidence the simple stellar model we introduced in this paper to study the shear statistics and extract useful information from future surveys. In particular, our analysis shows how the deviations from Gaussianity in the density field built by the non-linear gravitational dynamics translate into the PDF of the shear. This can also be seen in the Figures where we drew Gaussian PDFs for comparison. Hence we may hope to derive from shear maps some key properties of the structure of the density field (e.g., the amplitude of the many-body correlations and the shape of the PDF of the smoothed density contrast) which could shed some light onto the non-linear regime of gravitational clustering. In this respect, note that the separation of the shear, $\gamma$, into the factor $\kappa_{\text{min}}$ and the “normalized shear,” $\bar{\gamma}$, also provides a good separation between the dependence on the cosmological parameters and the projection effects (e.g., $\Omega_m, \Omega_\Lambda$ and $z_s$) versus the amplitude and the non-Gaussianities of the fluctuations of the density field which characterize the collisionless gravitational dynamics itself.

This is obvious for the convergence, $\kappa$, as recalled at the end of Section 3.2, but this feature also holds for the shear, $\gamma$.

However, in the near future, where the observational noise may still be important, one may first focus on the estimation of the cosmological parameters. To this order, our study is complementary to present analysis of the first few higher-order correlation functions of the shear field (Zaldarriaga & Scoccimarro, 2003, Schneider & Lombardi, 2003, Bernardineu, Mellier & van Waerbeke, 2002, and Takada & Jain 2003) based on results derived for cosmological spin-2 fields such as CMB polarization, where shape dependence and scale dependence of the shear is used to study non-Gaussianity induced by the underlying mass distribution. In any case, a few measures of non-Gaussianity (e.g., the skewness of the convergence, $\kappa$, see Pen et al., 2003) are useful to remove degeneracies between the various cosmological parameters (e.g., $\Omega_m$ vs $\sigma_8$) since a Gaussian is defined by only one quantity (its variance), while we would like to measure several quantities. Therefore, one may use our prediction for the PDF, $\mathcal{P}(\gamma)$, to build some useful new observables.

Before we can use our method to extract some information from weak gravitational lensing surveys, there are some points which need to be addressed. In particular, we should take into account the effects of a finite width for the distribution of sources over redshift, as in actual surveys. It would also be desirable to include realistic noise, such as that arising from the intrinsic ellipticity distribution of galaxies, and Poisson sampling due to the discreteness of the source distribution. The finite size of the catalogue is also worthy of consideration. These effects are beyond the scope of this paper, where we have mainly investigated the accuracy of our approach to describe the physics of weak gravitational lensing by large-scale structures. Having shown here that our simple stellar model provides a useful basis, we plan to investigate these points in detail in future work.

To incorporate noise in our calculations we need to convolve the pdf of shear which we have derived here with the noise pdf which is generally assumed Gaussian. A complete analysis is beyond the scope of the present paper. However, an order of magnitude estimate of the signal to noise ratio for the kurtosis of the shear (lowest order moment) has been given by Takada & Jain (2002). They conclude that for
reasonable values of survey parameters (rms intrinsic ellipticity \( \sigma_e = 0.4 \), number density of source galaxies \( n_{gal} = 30 \text{ arcmin}^{-2} \) and the survey area \( \Omega_{\text{survey}} = 25 \text{ degree}^2 \)) it is possible to achieve a high signal to noise ratio for the small angular scales which we have probed (although our cosmological models differ slightly with the model they have considered but such estimates are largely independent of the background cosmology). Since deviation from non-Gaussianity in the case of shear pdf is dominated by kurtosis one would expect roughly similar conclusions for the entire pdf too. However a detailed analysis in line with Munshi & Coles (2002) will be presented in a future work.

On the other hand, the present study already gives some information about delicate points which are of theoretical and practical interest. In particular, the underlying approximation in most analytical calculations (including ours) is the Born approximation. It neglects higher order correction terms in the photon propagation equation. The error introduced by such an approximation in the quasi-linear regime has been studied by several authors (e.g. Schneider et al. 1998). Although perturbative calculations tend to show that this error is negligible for lower order cumulants, clearly such an analysis is not possible in the highly non-linear regime. However, previous studies of convergence statistics and our present results regarding the shear PDF, which exhibit a good match between theoretical predictions and the simulation results, indicate that such corrections are negligible even in the highly non-linear regime.

With regard to numerical simulations, previous work for the predicted convergence statistics were made in comparison with ray-tracing experiments. In the present case, our analytical results for the shear, which are complimentary to the convergence results, are compared successfully with the numerical computations based on computing the full 3-d shear along every line of sight. The detailed match of our various predictions not only verifies our analytical calculations but also increases our confidence in the two different numerical methods.

ACKNOWLEDGEMENTS

This work has been supported by PPARC and the numerical work carried out with facilities provided by the University of Sussex. AJB was supported in part by the Leverhulme Trust. The original code for the 3-d shear computations was written by Hugh Couchman of McMaster University. There were many useful discussions with Antonio da Silva. DM acknowledges the support from PPARC of grant R328936. DM also acknowledges support from the University of Sussex during a visit where part of the work was completed. It is a pleasure for DM to acknowledge many fruitful discussions with members of Cambridge Leverhulme Quantitative Cosmology Group, including Jerry Ostriker and Alexandre Refregier.

REFERENCES

Babul A., Lee M.H., 1991, MNRAS, 250, 407
Barber A. J., Taylor A. N., 2002, submitted to MNRAS, (astro-ph/0212378)
Barber A. J., MNRAS, 2002, 335, 909
Bacon D.J., Refregier A., Ellis R.S., 2000, MNRAS, 318, 625
Balian R., Schaeffer R., 1989, A&A, 220, 1
Bartelmann M., Schneider P., 2001, Physics Reports, 340, 291B
Bernardeau F., 1994, A&A, 291, 697
Bernardeau F., Schaeffer R., 1992, A&A, 255, 1
Bernardeau F., Valageas P., 2000, A&A, 364, 1
Bernardeau F., Van Waerbeke L., Mellier Y., 1997, A&A, 322, 1
Bernardeau F., Mellier Y., Van Waerbeke L., 2002, A&A, 389, L28
Blandford R.D., Saust A.B., Brainerd T.G., Villumsen J.V., 1991, MNRAS, 251, 600
Couchman H. M. P., Barber A. J., Thomas P. A., 1999, MNRAS, 308, 180
Couchman H. M. P., Thomas P. A., Pearce F. R., 1995, Ap. J., 452, 797
Davis M., Peebles P.J.E., 1977, ApJS, 34, 425
Fry J.N., 1984, ApJ, 279, 499
Fry J.N., Peebles P.J.E., 1978, ApJ, 221, 19
Groth E., Peebles P.J.E., 1977, ApJ, 217, 385
Gunn J.E., 1967, ApJ, 147, 61
Hamilton A.J.S., Kumar P., Lu E., Matthews A., 1991, ApJ, 374, L1
Hoekstra H., Yee H. K. C., Gladders M. D., 2002, ApJ, 575, 995
Hui L., 1999, ApJ, 519, 622
Jain B., Seljak U., 1997, ApJ, 484, 560
Jain B., Seljak U., White S.D.M., 2000, ApJ, 530, 547
Jannink G., Des Cloiseaux J., 1987, Les polyn`eres en solution, Les editions de physique, Les Ulis, France
Jaroszynski M., Park C., Paczynski B., Gott J.R., 1999, ApJ, 365, 22
Jaroszynski M., 1991, MNRAS, 249, 430
Kaiser N., 1992, ApJ, 388, 272
Kaiser N., 1998, ApJ, 498, 26
Limber D.N., 1954, ApJ, 119, 665
Lee M.H., Paczynski B., 1990, ApJ, 357, 32
Miralda-Escudé J., 1991, ApJ, 380, 1
Munshi D., 2000, MNRAS, 318, 145
Munshi D., Coles P., 2002, MNRAS, 329, 797
Munshi D., Coles P., 2003, MNRAS, 338, 846
Munshi D., Jain B., 2000, MNRAS, 318, 109
Munshi D., Jain B., 2001, MNRAS, 322, 107
Munshi D., Wang Y., 2003, ApJ, 583, 566
Peebles P.J.E., 1980, The large scale structure of the universe, Princeton University Press
Munshi D., Bernardeau F., Melott A.L., Schaeffer R., 1999a, MNRAS, 303, 433
Munshi D., Coles P., Melott A.L., 1999b, MNRAS, 307, 387
Munshi D., Coles P., Melott A.L., 1999c, MNRAS, 310, 892
Munshi D., Melott A.L., Coles P., 1999d, MNRAS, 311, 149
Peacock J.A., Dodds S.J., 1994, MNRAS, 267, 1020
Pen U.-L., et al., 2003, ApJ, 592, 664
Premadi P., Martel H., Matzner R., 1998, ApJ, 493, 10
Schaeffer R., 1984, A&A, 134, L15
Schneider P., Ehlers J., Falco E. E., 1992, ‘Gravitational Lenses,’ Springer-Verlag, ISBN 0-387-97070-3
Schneider P., Weiss A., 1988, ApJ, 330, 1
Schneider P., Lombardi M., 2003, A&A, 397, 809
Schneider P., Van Waerbeke L., Jain B., Kruse G., 1998, MNRAS, 296, 873, 873
Scoccimarro R., Colombi S., Fry J.N., Frieman J.A., Hivon E., Melott A.L., 1998, ApJ, 496, 586
Scoccimarro R., Frieman J.A., 1999, ApJ, 520, 35
Scoccimarro R., Zaldarriaga M., 2003, ApJ, 584, 559
Szapudi I., Szalay A.S., 1993, ApJ, 408, 43
Szapudi I., Szalay A.S., 1997, ApJ, 481, L1
Takada M., Jain B., 2002, MNRAS, 337, 875
Takada M., Jain B., 2003, ApJ, 583, L49
Valageas P., Schaeffer R., 1997, A&A, 328, 435
Valageas P., 2000a, A&A, 354, 767
Valageas P., 2000b, A&A, 356, 771
Valageas P., 2002a, A&A, 382, 412
Valageas P., 2002b, A&A, 382, 477
Van Waerbeke L., Bernardeau F., Mellier Y., 1999, A&A, 342, 15
Van Waerbeke L., Mellier Y., Pelló R., Pen U.-L., McCracken H. J., Jain B., 2002, A&A, 393, 369
Van Waerbeke L., et al., 2000, A&A, 358, 30
Villumsen J.V., 1996, MNRAS, 281, 369
Wambsganss J., Cen R., Ostriker J.P., 1998, ApJ, 494, 298
Wambsganss J., Cen R., Xu G., Ostriker J.P., 1997, ApJ, 494, 29
Wambsganss J., Cen R., Ostriker J.P., Turner E.L., 1995, Science, 268, 274
Wang Y., Holz D. E., Munshi D., 2002, ApJ, 572L, 15
White M., Hu W., 2000, ApJ, 537, 1
Zaldarriaga M., Scoccimarro R., 2003 ApJ, 584, 559