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Influence of tidal dissipation on outcomes of binary–single encounters between stars and black holes in stellar clusters

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ABSTRACT
In the cores of dense stellar clusters, close gravitational encounters between binary and single stars can frequently occur. Using the TSUNAMI code, we computed the outcome of a large number of binary–single interactions involving two black holes (BHs) and a star to check how the inclusion of orbital energy losses due to tidal dissipation can change the outcome of these chaotic interactions. Each interaction was first simulated without any dissipative processes and then we systematically added orbital energy losses due to gravitational wave emission [using post-Newtonian (PN) corrections] and dynamical tides and recomputed the interactions. We find that the inclusion of tides increases the number of BH–star mergers by up to 75 per cent; however, it does not affect the number of BH–BH mergers. These results highlight the importance of including orbital energy dissipation due to dynamical tides during few-body encounters and evolution of close binary systems within stellar cluster simulations. Consistent with previous studies, we find that the inclusion of PN terms increases the number of BH–BH mergers during binary–single encounters. However, BH–star mergers are largely unaffected by the inclusion of these terms.

Key words: galaxies: star clusters: general – gravitation – methods: numerical – (stars:) binaries (including multiple): close – stars: kinematics and dynamics.

1 INTRODUCTION
Binary–single scatterings can frequently occur in the cores of dense globular clusters. These strong dynamical interactions result in cluster heating and are an important formation channel for stellar exotica and compact object binary systems (McMillan 1986; Davies, Benz & Hills 1993; Sigurdsson & Phinney 1993). There is a need to properly treat and determine the outcome of these interactions within stellar cluster evolution simulations. However, the task of analytically predicting the motions of three bodies in a closed system has been unsolved for hundreds of years (Poincaré 1892). The problem stems from the lack of a closed-form solution to the equations of motion. In the absence of an analytic solution to this problem, several recent works have tried to statistically describe and predict the outcome of these interactions (Stone & Leigh 2019; Ginat & Perets 2021; Kol 2021; Manwadkar et al. 2021; Parischewsky et al. 2021). Commonly used direct N-body and Monte Carlo simulation codes for evolving stellar clusters numerically integrate the equations of motions to determine the outcome of these chaotic interactions.

Dense stellar clusters could retain a sizeable fraction of stellar-mass black holes (BHs) that form when massive stars evolve and end their lives. (Morscher et al. 2015; Askar, Arca Sedda & Giersz 2018). Due to dynamical friction, these BHs can segregate to the centre of the stellar cluster where they can interact with each other and surrounding stars (Arca Sedda, Askar & Giersz 2018; Kremer et al. 2019a). This can result in numerous binary–single encounters involving multiple BHs and/or stars in the cores of stellar clusters. These dynamical interactions can lead to the formation, hardening, and ejections of binary BHs that could potentially merge due to gravitational wave radiation (e.g. Portegies Zwart & McMillan 2000; Moody & Sigurdsson 2009; Downing et al. 2010; Tanikawa 2013; Bae, Kim & Lee 2014; Ziosi et al. 2014). Therefore, the dynamical production of merging binary BHs in stellar clusters (e.g. Rodriguez, Chatterjee & Rasio 2016; Askar et al. 2017; Di Carlo et al. 2019; Banerjee 2021; Rodriguez et al. 2022) could be one of the channels for forming gravitational wave sources being observed with the ground-based LIGO/Virgo detectors (Abbott et al. 2016; 2021; The LIGO Scientific Collaboration 2021).

Recent studies examining merging BHs in dense environments have shown the importance of including orbital energy and angular momentum dissipation through gravitational waves during single–single (Samsing et al. 2020), binary–single (Samsing, MacLeod & Ramirez-Ruiz 2014; Samsing 2018), and binary–binary (Zevin et al. 2019; Arca Sedda, Li & Kocsis 2021) encounters in stellar clusters. The inclusion of these dissipative effects increases the number of merging binary BHs in stellar clusters and can also lead to eccentric BH mergers during few-body encounters involving BHs.

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(Banerjee 2018; Rodriguez et al. 2018a, b; Samsung, Askar & Giersz 2018c). Recently Samsung et al. (2022) have shown that dynamical single and binary BH encounters in discs of active galactic nuclei can also produce eccentric BH mergers. The detection of these eccentric mergers would be useful in constraining the astrophysical origin of binary BHs (Samsing & Ramirez-Ruiz 2017; Samsung et al. 2018a; Romero-Shaw et al. 2020; Zevin et al. 2021). In order to include dissipative effects of gravitational wave emission in the numerical computation of few-body interactions, post-Newtonian (PN) correction terms (e.g. Fackerell 1968; Blanchet 2014) to the Newtonian equations of motion are used (Kupi, Amaro-Seoane & Spurzem 2006; Harfst et al. 2008; Mikkola & Merritt 2008).

Scattering codes such as FEWBODY by (Fregeau et al. 2004) and TSUNAMI (Trani et al. 2016; Trani, Fuji & Spera 2019) have been developed to compute the outcome of dynamical interactions for a small number of objects. These integrators solve the N-body equations of motion to advance the system forward in time. The FEWBODY code has been used to compute the outcome of binary–single and binary–binary interactions within Monte Carlo simulation codes for stellar clusters such as MONTE CARLO CLUSTER SIMULATOR (MOCCA, Hypeki & Giersz 2013) and CMC (Rodriguez et al. 2022). There has been a lot of focus on including PN correction terms (up to 2.5PN) to capture the dissipative effects of gravitational wave emission when computing few-body interactions involving multiple BHs within stellar cluster simulations (e.g. Banerjee 2018; Rodriguez et al. 2018a, b). However, few studies have considered the effect of dissipation of tidal energy in few-body interactions involving BHs and stars (Ginat & Perets 2021).

Dynamical interactions between BHs and stars can frequently occur within stellar clusters. Such interactions can lead to the tidal disruption of stars (Fabian, Pringle & Rees 1975; Fragione et al. 2019; Kremer et al. 2019b, 2022). Binary–single interactions involving two BHs and a star are particularly interesting since the tidal disruption of stars by binary BHs may influence the properties of merging BHs (Lopez et al. 2019). Given that the inclusion of PN corrections to the equations of motion when computing the outcome of binary–single encounters between BHs leads to an increase in their merger rate. In this paper, we investigate how the inclusion of tidal dissipation can influence the outcome of binary–single encounters involving two BHs and a star. For this purpose, we use the TSUNAMI code (see Section 2.1) to carry out a few 100000 scattering experiments to systematically check how the inclusion of PN terms and tides can affect the outcome of binary–single encounters between two BHs and a star. In order to do this, each encounter was first computed without the inclusion of any dissipative effects, we then computed the same encounter with the inclusion of PN corrections. After this the encounter was recomputed with inclusion of tides and then finally with both PN corrections and tides. We find that the inclusion of tidal dissipation results in a significant increase of BH–star mergers. We show the statistical significance of this result and point out specific examples where the inclusion of tidal effects results in a different outcome. These results point out towards the importance of including the effects of tidal dissipation during few-body encounters in evolving stellar clusters.

In Section 2, we describe the TSUNAMI code and its main features as well as a brief explanation of how the interactions are set up. In Section 3, we then describe the initial set-up and results of the few-body encounters that we used to test the TSUNAMI code. In Section 4, we describe and show results from the initial data set for binary–single encounters that we took from the MOCCA-Survey Database I (Askar et al. 2017). The section also contains specific examples which highlight the importance of including tidal dissipation effects during these encounters.

2 METHODS

2.1 The TSUNAMI few-body code

We use the direct few-body integrator TSUNAMI (Trani et al. 2016, 2019) to integrate binary–single interactions. TSUNAMI employs the so-called algorithmic regularization chain scheme (e.g. Mikkola & Aarseth 1993; Mikkola & Tonikawa 1999; Hellström & Mikkola 2010). This scheme is a combination of three numerical techniques: Bulirsch–Stoer extrapolation (Bulirsch & Stoer 1964), logarithmic Hamiltonian regularization (Mikkola & Tonikawa 1999; Pretto & Tremaine 1999), and chain-coordinate system (Aarseth & Zare 1974; Heggie 1974; Zare 1974).

TSUNAMI also includes velocity-dependent corrections to the Newtonian equations of motion, including PN terms (e.g. Fackerell 1968; Blanchet 2014) of order 1, 2, and 2.5PN, and tidal forces (Press & Teukolsky 1977; Hut 1981; Samsing, Leigh & Trani 2018b). In the next section we describe in detail the model for the dynamical tide that is included in TSUNAMI.

2.2 Dynamical tide model

TSUNAMI uses the drag force model described in Samsing et al. (2018b). The drag force is given by

$$ F = -\delta \frac{v}{r^2} \times \frac{v}{v}, $$

(1)

where $n = 4$. $\delta$ is a normalization factor that can be estimated with

$$ \delta = \Delta E \times \frac{1}{2} \left[ a (1 - e^2) \right]^{-1/2} (GM)^{1/2} \mathcal{F}(e, n). $$

(2)

$\mathcal{F}(e, n)$ is the solution to the integral

$$ \mathcal{F}(e, n) = \int_{\theta_0}^{\pi} \frac{(1 + e \cos \theta) - (1 - e^2) / 2}{(1 + e \cos \theta)^{2n}} d\theta. $$

(3)

For any value on $n$, $\mathcal{F}$ can be written in a closed form. For $n = 4$, $\mathcal{F}$ evaluates to

$$ \mathcal{F}(e, n = 4) = \frac{\pi}{2} (2 + 7e^2 + e^4), $$

(4)

where $\theta_0$ is assumed to be $\pi$.

To obtain $\Delta E$ we consider a star with mass $m_1$ and radius $r_1$ on a parabolic orbit around a star with mass $m_2$, the energy loss due to tidal effects can be written as (Press & Teukolsky 1977b)

$$ \Delta E = \frac{G m_1^2 m_2}{m_1^2} \left( \frac{r_1}{r_2} \right)^6 T_2(\eta), $$

(5)

where $r_2$ is the pericentre distance of the orbit and $T_2$ is the tidal coupling constant of order 2. $T_2$, where $l > 2$, is much smaller than $T_2$ to the point where $T_2$ is one order of magnitude less than $T_2$, and for higher $l$ it is even smaller. Therefore they are excluded and only $T_2$ is included in the code. For a polytropic star, the tidal coupling constant can be found by a polynomial fit (Portegies Zwart & Meinen 1993):

$$ \log_{10} T_2(\eta) = A + Bx + Cx^2 + Dx^3 + Ex^4 + Fx^5, $$

(6)

where $x = \log_{10} \eta$ and the coefficients, $A, B, C, E$, etc., can be found in table 1 of Portegies Zwart & Meinen (1993). $\eta$ is calculated from
These equations were extended to hyperbolic and eccentric orbits by Mandel & Aarseth (2001), where they introduced a new parameter $\zeta$, defined by

$$\zeta = \eta \left( \frac{2}{1 + e} \right)^{1/2},$$

where $e$ is given by

$$e = 1 + \frac{\eta - 2}{2} \left( \frac{r_1^2}{r_2^2} \right)^{1/2}. \quad (9)$$

This $\zeta$ is then used instead of $\eta$ in equations (5) and (6).

The drag in force in TSUNAMI is implemented in the following way, given two objects, 1 and 2, where tides on 1 are excited by body 2:

(i) Calculate $\mathcal{F}$ analytically with equation (4).

(ii) Calculate the semimajor axis and eccentricity of the orbit, assuming that the orbit is a Keplerian orbit.

(iii) Use these values to calculate the energy loss using equation (5).

(iv) Estimate the drag force coefficient $\mathcal{F}$ using equation (2).

(v) Calculate the drag force vector between the two objects, $F_{1,2}$, using equation (1).

(vi) The acceleration due to tidal energy losses on object 1 caused by object 2 is then calculated as $a_{1,2} = F_{1,2}/m_1$.

This process is repeated for all objects in the system and the resulting acceleration is the corresponding sum of the vectors. This is done at every time-step for all non-compact objects.

2.3 Gravitational wave inspiral times

While TSUNAMI includes PN corrections, it is not numerically convenient to directly integrate the coalescence of an isolated binary. For this reason, when a BH–BH binary survives an encounter, we estimate its coalescence time analytically. An isolated BH binary will eventually merge due to energy losses through gravitational wave radiation emission. Peters (1964) presents a way of approximating this merger time:

$$t_{GW} = \frac{5}{304G^4} \frac{a^4}{m_1 m_2 (m_1 + m_2)} f(e), \quad (10)$$

where $c$ is the speed of light, $a$ is the semimajor axis, $e$ is the eccentricity, $G$ is the gravitational constant, and $m_1$ and $m_2$ are the masses of the two binary components. $f(e)$ is a factor that takes into account the binary eccentricity. As in Trani et al. (2022), we evaluate this factor numerically as:

$$f(e) = \frac{(1 - e^2)^{3/2}}{e} \frac{1}{\left(1 + \frac{2m_2}{m_1}\right) \left(1 + \frac{2m_1}{m_2}\right)^{1/2}} \int_0^e x^{1/2} \left(1 + \frac{12 + 11x^2}{36x^2} \right)^{1/2} dx. \quad (11)$$

Through interactions, such as a flyby, with other objects, the semimajor axis and eccentricity can change. By comparing the merger time before and after the interaction we can determine if, on a statistical scale, interactions with other objects increase or decrease this merger time. This is relevant since the code itself may not classify the interaction as a merger; however, the resulting binary may merge shortly after the interaction if the semimajor axis was reduced or the eccentricity increased as a result of the third object interacting with the binary.

2.4 Triple stability criterion

In order to classify a system as stable, TSUNAMI uses the approximate analytical criterion for dynamical stability from Mandel & Aarseth (2001). For a triple system with the objects 1, 2, and 3, we have an inner binary with $a_{in}$ and $e_{in}$ as well as an outer binary with $a_{out}$ and $e_{out}$. The mass ratio for the outer binary is $q_{out} = m_1 (m_2 + m_3)$, and we write the periastron separation as $R_{p_{out}}$. Mandel & Aarseth (2001) give the expression

$$R_{p_{out}} a_{in} = C \left[ (1 + q_{out}) + e_{out} \right]^{1/2}, \quad (12)$$

where $C = 2.8$ is given empirically. A triple configuration will be considered stable if $R_{p_{out}} > R_{p_{in}}$ and the formula holds for $q \leq 5$.

2.5 Initialization and outcomes of three-body encounters

All interactions carried out in this work are done in three-dimension (3D). FEWBODY is used to obtain the input data at $t = 0$ which is then provided to TSUNAMI in order to integrate the interactions. To generate the initial 3D positions and velocities coordinates at $t = 0$, we provided FEWBODY with the masses and radii of the three stars involved in the interaction, the semimajor axis and eccentricity of the binary, the impact parameter, and relative velocity of the interacting stars at infinity. The orbital elements of the interaction (inclination, longitude of the ascending node, argument of periastron, and time of periastron passage of the binary) were randomly selected by FEWBODY as described in section 3.1 of Fregeau et al. (2004). FEWBODY uses a seed to randomize these angles and for each interaction we select a seed from a large uniform distribution to ensure that the initial set-up is unique to each interaction. We present each data set we mention how many seeds are used for each interaction.

The MOCCA (Giersz et al. 2013; Hypki & Giersz 2013) is one of the most advanced codes used to simulated real size star clusters such as globular clusters. We use several different globular cluster simulations1 to extract binary–single interactions involving two BHs and a single main-sequence star where the star is either as the initial single or is initially a binary component. We will discuss these sets in more detail in Section 4.

Fig. 1 shows a schematic of outcomes during a binary–single interaction with two BHs and a star. The schematic is split into two sides, on the left side, the initial binary contains a star and a BH and the single is a BH. On the right side, the initial binary is a BH binary and the single is a star. The vertical lines represent the different outcomes that can occur: flyby, exchange, or a merger. There is also a possibility that the single breaks up the binary, an ionization, and the result is three singles. Another outcome not shown in the schematic is a stable hierarchical triple where we have an inner binary with the third object bound to this binary. Both of these outcomes are uncommon in our data set and are not included in the schematic. For all outcomes see Table 1.

We use notations similar to what is used in Fregeau et al. (2004) for bound systems and mergers. A bound system is represented by ‘l’, e.g. ‘BH1 [BH2: S]’ represents a binary with a BH and a star with an incoming BH. Mergers are represented by ‘+’, e.g. ‘BH1:BH2’

1Carried out as part of the MOCCA-Survey Database 1 (Askar et al. 2017).
represents a BH merger and ‘BH\(_1\):S’ represents a merger between the incoming BH and the star.

### 2.6 Binary mergers due to tidal dissipation

With the dynamical tides treatment in TSUNAMI, a few BH–star binaries with very high initial eccentricities would merge in an inspiral caused by the tidal effects before the single BH gets close. Since these mergers cannot be considered to be three-body interactions we need to filter them out. We do this by considering the binary to be a single body and we can thus work with analytical formulas for two-body interactions. We apply three criteria in order to find these ‘instant’ mergers:

(i) The merger occurs between the two initial binary components and one of the objects is a star.

(ii) The merger time returned by TSUNAMI is less than half of the time to pericentre for the orbit between the binary and the single.

(iii) The separation between the center of mass of the two merger objects and the third object is greater than five times the initial semimajor axis of the binary.

Any interaction which fulfils these three criteria is considered to be a merger not influenced by the single and filtered out and not included in our results reported in this paper.

### 2.7 Merger criterion

Mergers in TSUNAMI are classified using the ‘sticky star’ approximation. When the separation between two stars are less than the sum of their radii, TSUNAMI stops the interaction and records that a merger occurred during the interaction.
Table 2. Initial properties for the test set-up. For the semimajor axis we create eight bins equally distributed in log between 5 R⊙ and 1 au. For each semimajor axis bin, we calculate the maximum $b_{\text{max}}$ and pick 100 values from a uniform distribution between 0 and $b_{\text{max}}$.

| Property | Value |
|----------|-------|
| $M_*$   | 1 M⊙ |
| $M_{\text{BH}}$ | 10 M⊙ |
| $a$     | Min: 5 R⊙, max: 1 au |
| $e$     | 0 |
| $b$     | Min: 0, max: $b_{\text{max}}$ (see equation 13) |
| $V_\infty$ | 5 km s$^{-1}$ |

3 TEST SETUP

3.1 Initial properties

In order to see the effect of the different flags in a more controlled set we create a test set-up with initial properties as seen in Table 2. All stars in this set are Sun-like stars while the BHs have a mass of 10 M⊙. The semimajor axis is binned from a minimum of 5 R⊙ up to a maximum of 1 au with a total of eight uniformly distributed bins. For simplicity, we set the eccentricity to 0 and $V_\infty$ to 5 km s$^{-1}$. We are mostly interested in strong encounters and need to set our impact parameter accordingly. We do this by considering gravitational focusing in the limit where the velocity at pericentre is much higher than the velocity at infinity (Samsing, MacLeod & Ramirez-Ruiz 2014). This allows us to get a maximum value of the impact parameter using equation (13) where $G$ is the gravitational constant, $M_{\text{tot}}$ is the total mass of the three objects, $r_{\text{peri}}$ is the pericentre distance between the binary and the single, and $V_\infty$ is the velocity at infinity.

$$b_{\text{max}} = \sqrt{\frac{2G \cdot M_{\text{tot}} \cdot r_{\text{peri}}}{V_\infty^2}}. \quad \text{(13)}$$

We set the semimajor axis of the binary, $a$, as the pericentre between the binary and the single in order to get stronger interactions that are more likely to result in resonant flybys, exchanges, or mergers.

For each semimajor axis bin, we pick 100 impact parameters from a uniform distribution between 0 and $b_{\text{max}}$ and take the square root of these values. For each interaction we also assign three different seeds, this results in a total of 800 unique interactions and 2400 interactions in total. We create two similar data sets: one where we start with the star and a BH in the binary (test set-up 1) and one with the two BHs in the binary (test set-up 2). All initial parameters are the same for both sets and the only difference is whether the star is initially in the binary or as the single incoming object.

3.2 Results

The outcomes from the test set-up are shown in Tables 3, 4, and 5. Tables 3 and 4 show the outcomes when the initial binary contains a BH and a star (test set-up 1) while Table 5 shows the outcomes when the initial binary contains two BHs (test set-up 2). Some cells contain two rows, here the first number shows the fraction of interactions that ends with this outcome relative to the total number of interactions and the second number shows the total number of interactions with this outcome. The number inside the brackets indicates the number of unique interactions, i.e. identical interactions with the same outcome but with different seeds are only counted once. Looking at Table 3, we can see that for all outcomes; flybys, exchanges, and mergers, we see no significant differences between the no flags and PN terms run. However, when including tidal effects we see an increase in mergers and a drop in flybys and exchanges as a result of this. In Table 4, we take a closer look at the mergers from Table 3. As previously mentioned the number of mergers increase when tides are
involved. Since the cross-section for a merger with tides is higher than without, this is expected. Looking further into these mergers we can see that with no flags and only tidal effects we have no BH1:BH2 mergers; however, with PN terms and PN + tides we have one merger for each run. As expected, it is the number of BH–star mergers that is increased by tides which increase by approximately 54 per cent when tidal effects are included. We can split these mergers into intergers between the initial single incoming BH (BH1) and the initial binary BH (BH2). The increase in the number of BH2:S mergers (∼75 per cent) is noticeably higher than for BH1:S mergers (∼35 per cent) when comparing the number of mergers in the runs without tides to the runs with tidal effects. A reason behind this might be that the high-mass incoming BH strongly perturbs the orbit of the star to the point where it merges with the binary companion BH.

In Table 5, we show the outcomes for interactions where the initial binary contains two BHs (test set-up 2). In this set, we see no exchanges since it is generally unfavourable to exchange in an object with 10 times lower mass than any of the binary members. A majority of the interactions in all runs results in a flyby; however, we have more mergers (∼35 per cent increase) when including tides. We do not have any BH1:BH2 mergers so all of this increase is found in the BH:S mergers.

### 3.3 Example interaction

Here, we present an example of a BH1 [BH2 S] interaction from the test set-up where the inclusion of tides changes the outcome of the interaction. All initial parameters are constant in the test set-up except for the semimajor axis and impact parameter. The parameters are shown in Table 2 with semimajor axis $a = 0.117$ au and impact parameter $b = 8.6$ au.

The trajectories in the x–y plane for this interaction are shown in Fig. 2. Panel a shows the trajectory for an interaction with no additional processes included, for this run we initially have an exchange where we form a temporary triple with the inner binary consisting of the incoming BH and the star with the initial binary BH as the third outer object. The three objects quickly return to pericentre and we have a chaotic interaction which eventually ends with the formation of a BH–BH binary and the star as a single object. In panel b, we show the trajectory with tides included, in this case we have the same initial exchange and form a temporary bound triple with the incoming BH and the star as the inner binary. When the three objects return to pericentre we have a chaotic encounter which ends with a merger between the two initial binary components.

The resulting binary in the no flags run has properties $a = 0.6552$ au, $e = 0.905$, and $M_1 = M_2 = 10 M_\odot$. Using equation (10), we can get an estimate for the merger time of this binary due to gravitational wave radiation: $t_{GW} \approx 7.08 \times 10^5$ yr. For the tides run we use the sticky-star approximation to get the state vector of the centre of mass at the point of the merger. We assume no mass-loss and conserve linear momentum during the merger, and use the centre of mass as the remaining object. Calculating the orbital parameters for the orbit between the merger product and the third object we get $a = 0.953$ au and $e = 0.9662$, with the masses $M_1 = 10 M_\odot$ and $M_2 = 11 M_\odot$. This gives a merger time of $t_{GW} \approx 1.10 \times 10^6$ yr. From this we can see that, for this example, including tides not only leads to a BH–star merger; however, also the final binary will have a lower merger time compared to the binary that is created through the exchange in the run with no additional processes.

### 4 DATA FROM MOCCA

#### 4.1 MOCCA set 1

We extract approximately 82,000 binary–single interactions with two BHs and a star from several different MOCCA simulations of globular cluster evolution. These interactions are simulated using TSUNAMI with a seed picked from a uniform random distribution which results in different initial orientations for all interactions. All of these interactions are simulated using one seed for each interaction.

#### 4.1.1 Initial properties

The initial parameters for MOCCA set 1 (M1) are presented in Fig. 3. Each panel is split into two cumulative distributions, one for interactions with an initial BH–BH binary (M1 set-up 1) and one for interactions with an initial BH–star binary (M1 set-up 2). From left to right, the first panel shows the semimajor axis distribution which ranges from $10^{-7}$ to $10^4$ au. Set-up 1 has, on average, higher semimajor axis than the set-up 2. This might be connected to the way binaries are formed in three-body interactions in MOCCA. The semimajor axis is proportional to three times the average kinetic energy and thus if the binary components are very massive the semimajor axis can be very large. We can also see that in set-up 1 there is a small bump at around 8 au. This bump is related to primordial BH–BH binaries while most of the rest of the curve is dominated by dynamically formed BH binaries.

The eccentricity is shown in the next panel where set-up 1 has, on average, lower eccentricity than set-up 2. A reason for this might be that many of the BH–star binaries form dynamically and are frequently perturbed by surrounding single BHs and flyby interactions which can increase their eccentricity. We also have a few binaries with zero eccentricity. The BH–star binaries with zero eccentricity are most likely circularized by tidal forces (Hurley, Tout & Pols 2002) while for the BH–BH binaries, gravitational wave radiation can circularize the system (Peters 1964). The distribution of eccentricities in a relaxed system is expected to be thermal (Jeans 1919; Heggie 1975; Leigh et al. 2016). However, we can see that our distribution deviates from a thermal distribution and we have many higher eccentricity binaries. This is probably because such binaries are subject to frequent flyby interactions that can pump up their eccentricity. Additionally, the interactions that we extract originate from multiple clusters at different times during their evolution.

In the next panel, the impact parameter is shown, we can see that set-up 2 has, on average, lower impact parameter in relation to the semimajor axis compared to set-up 1. This is caused by the differences in semimajor axis. The first panel on the second row shows the initial velocity at infinity where we can see that set-up 1 has, on average, higher $v_\infty$ than set-up 2. In a relaxed system such as a globular cluster, we evolve towards energy equipartition. By assuming that the stars have masses proportional to the average mass in the cluster we can say that the average velocity for a star is close to the average velocity in the cluster. Since the average mass of a BH is larger than the average mass of regular, the average velocity of BHs is also lower. Thus, the average relative velocity between a star and a BH binary is larger than the one between a BH and a BH–star binary.

The middle panel on the second row shows the masses of stars for the interactions with both initial configurations; set-up 1 and 2. We can see that single stars have on average lower mass than the stars found in BH–star binaries in this set. A reason for this might be that it is difficult to have primordial binaries with very massive and
low-mass components. A binary between a massive BH and a star is also likely to be broken up by dynamical interactions. We can also assume that when a BH binary interacts with a star, the mass of the star follows the distribution of single stars in the system. However, stars in binaries have to be more massive than single stars since most BH–star binaries are formed in exchanges where usually the more massive star is bound to the BH. The final panel shows the BH masses and here the BH masses are very similar and the differences between the incoming BHs and the binary component BHs are very small.

The masses reflect the BH mass distribution connected with the assumed prescriptions for the evolution of massive stars in the MOCCA-Survey Database I (Askar et al. 2017) simulations. For binary/stellar evolution, these runs used prescriptions from Hurley et al. (2002). In some of the runs, the BH mass is modified according to the mass fallback prescriptions provided by Belczynski, Kalogera & Bulik (2002). The maximum BH mass produced in through the evolution of a single star was about 30 $M_\odot$. We can see from both distributions that there seems to be two groups of BHs, one group where the mass is between 10 and 15 $M_\odot$, which most likely corresponds to BHs without mass fallback and one group with mass $>20 M_\odot$, which most likely corresponds to BHs with mass fallback (Zhang, Woosley & Heger 2008). For the evolution of BH progenitors, these simulations did not consider metallicity-dependent winds, pair/pulsational pair instability supernovae, and the supernova explosion model (Fryer et al. 2012; Spera & Mapelli 2017; Woosley 2017; Vink & Sander 2021).

### 4.1.2 Results

Tables 6, 7, and 8 show the outcomes for set M1. Table 6 shows all outcomes for interactions from M1 set-up 1, we can see that a majority of interactions result in a flyby followed by exchanges. The number of mergers is small; however, we can see that the inclusion of tides increases the merger rate by approximately 40 per cent. We
also have a few ionizations where the incoming single causes the binary to ionize and the end result is three single objects. In Table 7, we take a closer look at the number of mergers. Without PN terms we do not see any BH–BH mergers; however, with PN terms included we see three for the run with only PN terms and two for the run with PN terms and tides. The inclusion of tides leads to an increase of approximately 40 per cent for the BH:S mergers. The merger rates between BH1 and the star increase by approximately 100 per cent while the merger rates between the two initial binary components increase by approximately 26 per cent. Compared to the test set-up (see Section 3), the tidal effects increase the number of BH1:star mergers to a much larger extent in M1; while for the test set-up, the number of BH2:star mergers see the largest increase. The reason for this might be that the binaries found in MOCCA are more stable and require stronger perturbations for the two binary components to merge. The binaries from MOCCA have also, on average, higher semimajor axis than the binaries we set up in the test set-up. The masses are also very different; in the test set-up the BH masses are always 10 times higher than the star mass. In the MOCCA set, this varies from interaction to interaction.

In Table 8, we show the number of outcomes for the interactions from M1 set-up 2. In this data set, a large majority of the outcomes results in flybys and a small fraction as exchanges. The number of mergers is low compared to the total number of interactions; however, if we compare the runs with and without tides, we can see that the inclusion of tides increases the number of BH:S mergers by approximately 140 per cent. BH:BH mergers are only found in the two runs with PN terms and the large increase in mergers is due to an increase in BH:S mergers. We also have a very small amount of ionizations; however, this number is consistent for all four runs. With PN terms we see one stable hierarchical triple where the BH binary captures the star in a wide orbit around the binary. TSUNAMI uses the triple stable criterion described in Section 2.4 in order to classify hierarchical systems as stable. The fact that we only see one of this outcome means that, in our data set and with this classification criteria, it is very unlikely to form stable hierarchical triples.

### 4.2 MOCCA set 2

#### 4.2.1 Initial parameters

A large majority of the interactions in M1 ends as a flyby and the interactions are quite distant, weak, and non-resonant. In order to get a better understanding of how the additional processes affects stronger interactions we create a filtered MOCCA set called MOCCA set 2 (M2). This set is created by integrating all interactions in M1 without any additional processes and a selected seed such that the interaction is co-planar. From these runs, all non-flybys are extracted and assigned five seeds for each interaction which results in a set of 5600 unique interactions and 28000 total interactions. We split the interactions into two groups; interactions where the star is initially in the binary and the incoming object is a BH (BH [BH:S], set-up 1) and interactions where the star is initially the incoming object and the binary is a BH–BH binary (S [BH BH], set-up 2).

The initial parameters are shown in Fig. 4. The initial semimajor axis distribution is found in the first panel on the upper row. We can see that on average, set-up 2 have a lower semimajor axis than set-up 1. We have two groups in set-up 2, as discussed in M1, this is related to primordial and dynamically formed binaries. Primordial binaries have, on average, lower semimajor axis and correspond to the left-hand side of the curve while dynamically formed binaries correspond to the right-hand side. The second panel shows the eccentricity distribution. The impact parameter distribution is shown in last panel of the upper row. The first panel on the lower row shows the velocity at infinity distribution. The second panel on the lower row shows the star mass and the final panel on the lower row shows the BH mass. Since this is a subset of M1, the initial parameter distributions are very similar, for explanations and discussions related to the initial parameters see Section 4.1.1.

#### 4.2.2 Results

This section will present the results from set M2 where we split up the interactions initially containing a BH–star binary interacting with a single BH (M2 set-up 1), shown in Tables 9 and 10, and interactions initially containing a BH–BH binary interacting with a single star (M2 set-up 2), shown in Table 11.

We start by looking at the outcomes for interactions with a BH–star binary (M2 set-up 1) in Table 9. A large majority of these interactions results in either a flyby, an exchange, or an ionization. The number of flybys and exchanges decreases when tides are included; however, the number of ionizations is consistent among all four runs. The decrease in flybys and exchanges with tides included are due to the increase in mergers. Without tides we have 103 mergers in the no flags run and 100 mergers for the PN terms run while we have 175 mergers for both the run with tides and the run with PN terms and tides. This increase is significant; however, we have to consider that each interaction is run five times with different seeds. Thus we can look at the numbers in bracket and see that the increase in mergers is approximately 29 per cent which is more in line with the other...

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### Table 6. M1 set-up 1 outcomes: The first column shows which flags were used and the second column shows the total number of completed interactions. The remaining columns show the number of flybys, exchanges, mergers, and ionizations, respectively. The rows have the same description as provided in the caption of Table 3.

|      | Total | Flybys | Exchanges | Mergers | Ionizations |
|------|-------|--------|-----------|---------|-------------|
| No flags | 5148  | 0.645  | 0.269     | 0.018   | 0.068       |
|       | 3321  | 1382   | 94        | 102     | 351         |
| PN    | 5143  | 0.643  | 0.269     | 0.020   | 0.068       |
|       | 3309  | 1381   | 102       | 351     |             |
| Tides | 5057  | 0.647  | 0.257     | 0.028   | 0.068       |
|       | 3270  | 1302   | 140       | 345     |             |
| PN + tides | 5051  | 0.645  | 0.255     | 0.031   | 0.069       |
|       | 3260  | 1289   | 156       | 346     |             |

### Table 7. M1 set-up 1 mergers: The first column shows which flags were used and the second column shows the total number of mergers. The rest of the columns show the number of BH–BH mergers, mergers between the incoming BH and the star, and the mergers between the initial binary components. The rows have the same description as provided in the caption of Table 3.

|      | Mergers | BH1:BH2 | BH1:S | BH2:S |
|------|---------|---------|-------|-------|
| No flags | 94      | 0       | 0.351 | 0.649 |
|       |        | 0       | 33    | 61    |
| PN    | 102     | 0.029   | 0.314 | 0.657 |
|       |        | 3       | 32    | 67    |
| Tides | 140     | 0       | 0.450 | 0.55  |
|       |        | 0       | 63    | 77    |
| PN + tides | 156     | 0.013   | 0.487 | 0.50  |
|       |        | 2       | 76    | 78    |
two sets. By looking at these mergers in more depth in Table 10, we can see that the number of BH:BH mergers increases very slightly in the two runs where we use PN terms compared to the two runs where we do not. For the runs without PN terms we have five BH:BH mergers; however, looking at the number of unique mergers we only have one. The binary is very tight at $a = 2.7 \times 10^{-2}$ au and the impact parameter is very low at $b = 8.59 \times 10^{-3}$ au. The single BH is travelling almost head-on and colliding very quickly with the binary component BH. Since the binary component has much higher mass than its binary companion star, the BH is always very close to

Table 8. M1 set-up 2 outcomes: The first column shows which flags were used, the second column shows the total number of complete interactions, and the third column shows the number of flybys. The fourth column shows the number of exchanges and the fifth column shows the total number of mergers. The sixth and seventh columns show the number of BH–BH mergers and BH–star mergers, respectively. The eighth column shows the number of ionizations and the ninth and final column shows the number of stable triples. The rows have the same description as provided in the caption of Table 3.

| Total | Flybys | Exchanges | Mergers | BH$_1$:BH$_2$ | BH$_1$:$S$ | Ionizations | Stable triples |
|-------|--------|-----------|---------|---------------|-------------|-------------|----------------|
| No flags | 76 842 | 0.9897 | 3e-4 | 0.01 | 0 | 0.01 | 9e-5 | 0 |
| PN | 76 838 | 0.9898 | 4.1e-4 | 0.01 | 2.6e-5 | 0.01 | 9e-5 | 1.3e-5 |
| Tides | 76 833 | 0.976 | 3e-4 | 0.023 | 0 | 0.023 | 9e-5 | 0 |
| PN + tides | 76 818 | 0.976 | 3.9e-4 | 0.023 | 2.6e-5 | 0.023 | 9e-5 | 0 |

Figure 4. Cumulative distribution plots for initial properties distributions in the M2 set (from left to right); semimajor axis, eccentricity, impact parameter, velocity at infinity, star mass, and BH mass.

Table 9. M2 set-up 1 outcomes: For a description of the columns, see Table 6. The rows have the same description as provided in the caption of Table 3.

| Total | Flybys | Exchanges | Mergers | Ionizations |
|-------|--------|-----------|---------|-------------|
| No flags | 9703 (1945) | 0.280 | 0.605 | 0.0106 | 0.105 |
| PN | 9693 (1945) | 0.280 | 0.604 | 0.0103 | 0.105 |
| Tides | 9286 (1946) | 0.263 | 0.610 | 0.0188 | 0.108 |
| PN + tides | 9287 (1946) | 0.263 | 0.610 | 0.0188 | 0.108 |
the centre of mass and for all seeds that we have investigated this interaction leads to a merger. The two runs with PN terms have the same five BH:BH mergers as previously discussed; however, there is also one additional interaction where the two BHs merge for one seed only.

The big increase with tides is, as expected, found in BH:S mergers where the increase in BH$_1$:S mergers is approximately 90 percent both when considering the total and unique number of mergers. The increase in BH$_2$:S mergers is a bit less at approximately 71 percent for the number of total mergers and 33 percent for the number of unique mergers. We find that a few of these BH:S mergers happen shortly after the first close encounter with the single, this is most likely due to energy being transferred from the binary to the single which shrinks the binary to the point where the tidal effects cause an inspiral. These binaries that inspiral shortly after the interaction may explain the rest of the increase in number of mergers when tides are included.

The outcomes for M2 set-up 2 are shown in Table 11. The majority of the interactions end as flybys, we have very few exchanges and the number of ionizations is constant for all four runs. The inclusion of tides increases the total number of mergers by approximately 124 percent and the number of unique mergers by 109 percent. The total number of BH:BH mergers increases by approximately 100 percent when PN terms are included; however, the number of mergers is low and thus this may be due to statistical fluctuations. The change in the number of unique mergers is more difficult to quantify since the numbers are very low (two or three); however, from the results we have, we get one less unique merger with PN terms compared to no flags but one more merger with PN + tides compared to only tides. Thus, due to the extremely small sample size, it is not possible to draw any conclusions for this and the differences we see may be due to statistical fluctuations.

The total number of BH:S mergers increase by approximately 125 percent when including tides and the unique number of mergers increase by approximately 110 percent. Compared to the other sets, this is the largest increase in the number of BH:S mergers we have found. However, this is a subset of M1 where we have taken out the most resonant and hard encounters thus we should expect to find a larger increase in merger rates when dissipative terms are included.

4.3 Example interaction

Here, we introduce an example interaction where the change of flags changes the outcome of the encounter. The initial parameters are found in Table 12 and the trajectories for the runs are found in Fig. 5. For the no flags and PN runs (panel a and b) this interaction results in a flyby and the differences between the two runs are insignificant. When tides (panel c) are included, the interaction is more interesting; when the binary and single gets close, we form a new binary consisting of the star and the incoming BH with the third BH bound to this newly formed binary. After some time the binary and single have another close encounter where the star is exchanged out of the binary and kicked out of the interaction. The result is a BH–BH binary and an unbound star.

With PN + tides included the start of the encounter is very similar to the run with only tides; the star and the incoming BH forms a new binary with the other BH bound to them. However, at the second close passage we have another exchange of the star. The star is exchanged back to the initial binary BH and shortly after inspirals and merges with the BH.

In both the tides and the PN + tides runs we form a BH–BH binary after the interaction, we can thus look at the merger times due to gravitational wave radiation. We use equation (10) for this and find that for the tides run we have $a = 8.15$ au, $e = 0.79$, $M_1 = 20.36 M_\odot$, and $M_2 = 19.88 M_\odot$ which gives us $\tau_{GW} \approx 2.42 \times 10^{14}$ yr. For the PN + tides run we use the sticky-star approximation to merge the star and the BH, we do not account for any mass-loss and use the centre of mass at the final time-step as the new merger product. By looking at the orbit between the merger product and the third object we find that the orbit is bound with $a = 6.27$ au and $e = 0.81$. The mass of the merger product is $21.01 M_\odot$ while the mass of the other BH is $20.36 M_\odot$. This gives a merger time of $\tau_{GW} \approx 5.82 \times 10^{13}$ yr. The merger times of both of these binaries are very long, longer than a Hubble time. However, what we can see is that the binary that is created after the merger has a longer merger time. This binary can then encounter additional objects in the cluster and harden to the point where the BHs will merge.

5 SUMMARY AND CONCLUSIONS

We have performed binary–single scattering experiments using the TSUNAMI code with both manually set up interactions as well as interactions extracted from the MOCCA-SURVEY Database I. Each interaction was simulated four times using: (i) pure N-body, (ii) with energy dissipation due to 2.5PN terms, (iii) with energy dissipation due to tidal forces, and (iv) with energy dissipation due to both PN terms and tidal forces.

We have found that the inclusion of both PN terms and tidal forces are important for binary–single interactions involving two BHs and a single star. Fig. 6 shows the fraction of BH:S mergers for both set-ups of our three data sets relative to the total number of interactions for each set. These histograms show that the inclusion of tides increases the merger rates between BHs and stars during binary–single interactions by a significant amount for both set-ups. The magnitude of the increase is, in our data sets, dependent on the initial configuration of the data sets and varies significantly among our three sets.

The inclusion of PN terms increases the number of BH:BH mergers but does not affect the number of BH:star mergers; however, inclusion of tidal forces increase the number of BH–star mergers but do not affect the number of BH–BH mergers in a significant way. The magnitude of the increase in both BH:star and BH:BH mergers from the inclusion of tides and PN terms depends on the initial parameters of the interaction.

The inclusion of PN terms increases the number of BH:BH mergers in all simulated data sets, in test set-up 1 we get one merger with PN terms and zero without them. The runs without PN terms do not result in any BH:BH mergers in M1 set-up 1 or 2, however, with PN terms included we see either two or three BH:BH mergers in each run.
Table 11. M2 set-up 2 outcomes: For a description about the columns, see Table 8, with the exception that we do not find any stable triples in this set-up and therefore do not include that column here. The rows have the same description as provided in the caption of Table 3.

| Total | Flybys | Exchanges | Mergers | BH₁BH₁ | BH₁S | Ionizations |
|-------|--------|-----------|---------|--------|------|-------------|
| No flags | 18 176 (3638) | 0.98 | 0.6e-4 | 0.188 | 1.65e-4 | 0.0187 | 5.5e-5 |
| 17 821 (3638) | 12 (12) | 342 (264) | 3 (3) | 339 (262) | 1 (1) |
| PN | 18 164 (3638) | 0.98 | 9.36e-4 | 0.188 | 3.30e-4 | 0.0187 | 5.5e-5 |
| 17 801 (3637) | 17 (15) | 342 (278) | 6 (2) | 339 (276) | 1 (1) |
| Tides | 18 152 (3638) | 0.957 | 7.17e-4 | 0.042 | 1.65e-4 | 0.042 | 5.5e-5 |
| 17 371 (3636) | 13 (12) | 767 (553) | 3 (2) | 764 (551) | 1 (1) |
| PN + tides | 18 158 (3638) | 0.956 | 7.16e-4 | 0.0436 | 3.86e-4 | 0.043 | 5.5e-5 |
| 17 353 (3635) | 13 (12) | 791 (568) | 7 (3) | 784 (565) | 1 (1) |

Table 12. Initial properties for example interaction extracted from M2 set-up 1.

| Property | Value |
|----------|-------|
| M_{BH₁} | 20.36 M_{⊙} |
| M_{BH₂} | 19.88 M_{⊙} |
| M_{*} | 1.14 M_{⊙} |
| a | 1.13 au |
| e | 0.98 |
| b | 14.01 au |
| V_{∞} | 26.58 km s⁻¹ |

When including tides in the case where we start with the star in the binary (set-up 1) we see an increase in both the number of mergers between the incoming single BH and the star as well as an increase in the number of mergers between the star and the binary companion BH. As discussed above, the magnitude of the increase is different for our data sets; in our test set-up, we find a 54 per cent increase in total number of mergers. For BH₁:S mergers we see an increase of 35 per cent and for BH₂:S mergers, an increase of 75 per cent. For M1 we see an increase in the number of total mergers of approximately 40 per cent while the number of BH₁:S mergers increase by approximately 100 per cent and BH₂:S mergers increase by approximately 26 per cent.

In M2, we can look at the total number of mergers as well as the number of unique mergers: With tides, we find an increase of approximately 70 per cent for the total number and 29 per cent for the number of unique mergers. The increase in the number of BH₁:S mergers is the same for the total number and number of unique mergers; approximately 90 per cent. The total number of B₂:S mergers increases by approximately 71 per cent and the number of unique mergers increases by 33 per cent. The reason behind the differences may be the masses of the objects and the initial set-up of the interactions. In the test set-up, we see a larger increase in the number of BH₁:S mergers compared to the other sets. This is because most the binaries in M1 and M2 are and not as tight as the ones we created in the test set-up; thus, it is harder for the incoming BH to perturb the orbit of the star to the point where it merges with the binary companion.

Tidal effects, in the case where the initial set-up is a BH–BH binary and a single star (set-up 2), increase the number of BH–star mergers; however, the magnitude of the increase is, similarly to the case previously discussed, dependent on initial set-up. For the test set-up, we see a 30 per cent increase in BH–star mergers with tidal effects compared to without. For M1, this increase is much larger at approximately 137 per cent. The increase is also larger for M2 at 125 per cent (total) and 110 per cent. Including PN terms in addition to tidal effects, slightly increases the number of BH–star mergers in the test set-up and M2 but decreases the number in M1.

The results summarized above show that there is a significant increase in the number of mergers between BHs and stars when tidal dissipation is included when computing the outcome of few-body interactions. This has important consequences on rates for transient events, like tidal disruptions caused by stellar-mass BHs within stellar clusters (Perets et al. 2016; Kremer et al. 2019b, 2021). Therefore, our findings indicate the importance of including tidal dissipation for correctly computing the outcome of close encounters within stellar cluster simulations.

The inclusion of PN terms in addition to tides seems to have different effects on the number of BH–star mergers depending on which data set we use; for the test set-up, we see less BH–star mergers in the PN + tides run compared to the tides run. In M1, we have slightly more mergers between the incoming BH and the star with both PN + tides but no significant change in the number of mergers between the two initial binary components. For M2, we see no difference in the number of BH–star mergers between the PN + tides and the tides run.

The number of ionizations remains constant and we cannot see any significant difference among our runs that would indicate a dependence on additional processes. For individual interactions, the inclusion of additional processes may change the outcome to or from an ionization; however, on a larger statistical scale there are no significant differences.

We found that the inclusion of the ‘impulsive’ description of dynamical tides causes a very quick merger for a non-negligible number of binaries, to the point where the single does not interact with it. This only affects binaries from the M1 and M2 data set that had a very high eccentricity and thus a low pericentre distance. Since the eccentricity is set to zero in the test set-up, we do not see any of these very fast mergers, even though we have a very low initial semimajor axis for a few binaries. For M1, we find that 1.56 per cent of all interactions merge before the three-body interaction, these quick mergers are found with the criteria described in Section 2.6. For M2, we find that in 4.1 per cent of all interactions, the binary mergers before interacting with the single; however, if we only consider unique interactions (i.e. filter out identical interactions with different seed), we find that 0.9 per cent of binaries merge without the influence from the third object. The initial set-up for M1 and M2 were taken from the MOCCA-Survey Database I simulations, in MOCCA, binary evolution is updated at the time of the interaction using prescriptions in the BSE code and then the FEWBODY code is called to compute the outcome of the interaction. Our results indicate that the treatment for tides in the BSE code does not capture tidally...
Figure 5. The x–y plane trajectory for the interaction with no flags (a) which results in a flyby, PN (b) which also results in a flyby, tides (c) which results in an exchange, and PN + tides (d) which results in a merger between the two initial binary components.

Figure 6. Histogram showing the fraction of BH:S mergers for both set-ups of the three data sets. The panels on the left show the BH$_1$:S and BH$_2$:S mergers for set-up 1 and the panels on the right show the BH$_{1,2}$:S mergers for set-up 2. The error bars show the Poisson error distribution as $\pm \sqrt{N}$. 
induced mergers of very eccentric binaries, whereas when we use the dynamical treatment of tides with TSUNAMI, we see that the binary components merge on very short time-scales. This points towards the need to update the treatment for tides in binary population synthesis codes like BSE, as they may be underestimating the merger times for binaries with small pericentre distance values.

Many BH:S mergers leave behind a weakly bound BH binary with a very high eccentricity and low pericentre distance. It is possible that due to GW radiation during the pericentre passage that the BHs merge or that the orbit shrinks rapidly into an inspiral. However, since the orbit is very wide, the time it would take for this binary to complete one orbit is high and it is possible that a third object may disrupt the binary before this happens. It is possible to compare the interaction time-scale in the clusters to the period of the binary to get an estimate of how many binaries would have time to complete one orbit before interacting again; however, estimating the interaction time-scale is very difficult since this is very dependent on the cluster density, velocity dispersion, and where in the cluster the particular interaction takes place. Therefore, this paper puts more focus on the merger rates during the interactions and leaves the possibility of following the binaries that are created as a result of the mergers for future studies.

Our results show an increase in mergers during binary–single interactions when additional processes are included. This increased merger rate may lead to a decrease in binary fraction which might influence overall cluster evolution. The inclusion of tidal effects is suspected to also have a large impact on close encounters when more than one star is involved. In interactions with two or more main sequence stars, the tidal effects may cause a merger between two main sequence stars which may in turn lead to a larger population of blue stragglers in the cluster. Additionally, tidal dissipation may also increase the number of disruption events involving interactions between BHs and evolved stars like giants (Ivanova et al. 2010, 2017) and white dwarfs (Rosswog, Ramirez-Ruiz & Hix 2009). This can be an important channel for dynamically forming compact binary systems in dense stellar clusters. We will study this in future publications in which we plan on carrying out stellar cluster simulations with the MOCCA code in which TSUNAMI (with tidal dissipation and PN corrections) will be used to compute the outcome of close encounters between two and more bodies.

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DATA AVAILABILITY

Input and output data for the scattering experiments carried out in this paper will be shared on request to the corresponding author. The TSUNAMI code and the simulated data will be shared on reasonable request to Alessandro A. Trani.

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