ABSTRACT

This paper reports X-ray spectral observations of a relatively nearby (z = 0.048) BL Lac object 1ES 1959+65, which is a potential TeV emitter. The observations include 31 short pointings made by the Unconventional Stellar Aspect (USA) experiment on board the Advanced Research and Global Observation Satellite (ARGOS), and 17 pointings by the Proportional Counter Array (PCA) on board the Rossi X-Ray Timing Explorer (RXTE). Most of these observations were spaced by less than 1 day. 1ES 1959+65 was detected by the ARGOS USA detector in the range 1–16 keV, and by the PCA in the 2–16 keV range but at different times. During the closely spaced RXTE observations beginning on 2000 July 28, the ending of one flare and the start of another are visible, associated with spectral changes, where the photon index $\Gamma$ ranges between $\sim 1.4$ and $\sim 1.7$, and the spectrum is harder when the source is brighter. This implies that 1ES 1959+65 is an X-ray-selected BL Lac type (XBL) blazar, with the X-ray emission likely to originate via the synchrotron process. The USA observations reveal another flare that peaked on 2000 November 14 and doubled the flux within a few days, again associated with spectral changes of the same form. The spectral variability correlated with the flux and timing characteristics of this object that are similar to those of other nearby BL Lac objects and suggest relativistic beaming with a Doppler factor $\delta \geq 1.6$ and magnetic fields on the order of a few milligauss. We also suggest that the steady component of the X-ray emission—present in this object as well as in other XBLs—may be due to the large-scale relativistic jet (such as measured by Chandra in many radio-loud active galactic nuclei) but pointing very closely to our line of sight.

Subject headings: BL Lacertae objects: individual (1ES 1959+65) — X-rays: galaxies

1. INTRODUCTION

Over a dozen BL Lac objects have been detected at GeV energies (Mukherjee et al. 1997), but only a few nearby BL Lac objects have been identified at TeV energies so far. Mrk 421 ($z = 0.031$; Punch et al. 1992) and Mrk 501 ($z = 0.034$; Quinn et al. 1996) are now strongly confirmed sources, seen by more than one ground-based atmospheric Cerenkov telescope (ACT) at or above the 5 $\sigma$ level. Two more, 1ES 2344+514 ($z = 0.044$; Catanese et al. 1998) and PKS 2155–304 ($z = 0.116$; Chadwick et al. 1999), have been detected only once and are less conclusive. This strongly suggests that low-redshift X-ray–selected BL Lac objects (XBLs) such as these may be the only extragalactic $\gamma$-ray sources observable at TeV energies. This is because, on the one hand, more distant objects would have their TeV emission strongly absorbed by the extragalactic background light (EBL), and on the other, bright X-ray emission, presumably originating from synchrotron radiation and thus revealing the distribution of radiating particles, implies even higher intrinsic GeV–TeV emission (see, e.g., Tavecchio, Maraschi, & Ghisellini 1998).

The BL Lac object 1ES 1959+65 ($z = 0.048$) is an XBL also present in the third EGRET catalog with an average measured flux of $1.8 \times 10^{-7}$ photons cm$^{-2}$ s$^{-1}$ for $E > 100$ MeV (Hartman et al. 1999). It is thus a natural source for TeV emission, and Stecker, De Jager, & Salamon (1996), using simple scaling arguments, have predicted for it the third highest flux above 0.3 and 1 TeV, after Mrk 421 and Mrk 501. More recently, Costamante & Ghisellini (2002) have also pointed it out as a candidate TeV emitter. Quoting Weeckes (1999), the Utah Seven Telescope Array has reported the detection of 1ES 1959+65 based on 57 hr of observation in 1998 (Kajino et al. 1999), with an energy threshold of 600 GeV. The flux level was not reported, but the total signal was at the 3.9 $\sigma$ level. This is not normally considered sufficient to claim the detection of a new source; however, within this database there were two epochs that were selected a posteriori that gave signals above the canonical 5 $\sigma$ level. This source has not yet been confirmed by any other group; it was observed by the Whipple group, but no flux was detected (Catanese et al. 1997). Of particular interest in this source is also the fact that its distance is of the same order as the two other confirmed TeV sources, making it a good candidate to probe the EBL that can, in turn, probe cosmological problems such as the formation of gal-
axies (see, e.g., Primack et al. 1999; Guy et al. 2000). 1ES 1959+65 is also part of a 200 mJy radio-selected sample at 5 GHz (Marcha et al. 1996; Bondi et al. 2001), and it was seen in the Einstein Slew Survey (Perlman et al. 1996). Multiple photometric optical values were found in the literature, showing a great variation of the source brightness in the optical band from $V = 16$ to $V = 12.8$. A complete study of the optical band can be found in Villata et al. (2000), where variability on short timescales (a few days) was reported.

The capabilities of the Unconventional Stellar Aspect (USA) and the Proportional Counter Array (PCA) instruments to monitor the X-ray emission are particularly well suited for the detailed study of the X-ray energy spectrum of 1ES 1959+65 and its temporal evolution. Here we present 31 USA and 17 PCA observations of 1ES 1959+65. The PCA data are two sets of intraday observations, obtained from unpublished Rossi X-Ray Timing Explorer (RXTE) archival data, that span a few days in which the end of a flare and the beginning of another flare are detected. The USA observations are short daily monitoring observations that span 7 weeks, ending with a strong flare, where the X-ray flux tripled in 7 days (the USA data were taken in the context of a multiwavelength campaign that was not successful in the TeV range because of bad weather at the ACT site). In §2, we report the details of the observations; in §3, we report the main features of the data; in §4, we discuss the implications of the X-ray spectral variability and the quiescent X-ray emission detected in 1ES 1959+65; we present our conclusions in §5.

2. OBSERVATIONS AND DATA ANALYSIS

2.1. USA Data

The USA Experiment observed 1ES 1959+65 about once per day between 2000 September 21 (MJD 51,808) and November 11 (MJD 51,859). (For a detailed description of the USA experiment, see Ray et al. 2000, Wood et al. 2000, and Shabad 2000.) The USA detector was used in the “ping” mode, where within one observation the source is pointed at 2 or 3 times for $T_{ON}$ seconds and then a background is taken for $T_{OFF}$ seconds, where typically $T_{ON} = T_{OFF} = 60$ s and an observation lasted for $\sim 300$ s (for more details on the “ping” observation mode, see Giebels et al. 2000). This type of observation has the advantage of not relying on a background model but rather on a direct measurement of the background, but results in some loss of signal within an observation period. USA data were extracted from FITS-formatted files using CFITSIO. Care was taken to use an OFF position as devoid as possible of bright X-ray sources and at least 2° away from the source. Observations were made in low-background regions of the orbit, where the counting rate was approximately 30 counts s$^{-1}$ or 1% of the Crab level in the same conditions. The segments of the observation where the background is too high, especially at the beginning and the ending of each observation, are rejected. For the present investigation, we use only observations where at least two pairs of ON-OFF pointings are available for each observation. Within those restrictions, 31 observations were used from which $\sim 16$ ks of data were selected.

We present the X-ray light curve for 1ES 1959+65 in Figure 1; in that figure, the standard deviation of the average in the background was added in quadrature to the error on the count rate for each observation. The data were then corrected for obscuration by the instrument support structure when necessary, and also for the collimator response. Every point in the light curve is a single observation. The light curve (normalized to the USA Crab rate in the energy range defined below, or 3500 counts s$^{-1}$) for the total range is shown in the top panel of Figure 1.

The data were taken in the spectral mode, where the instrument integrates a spectrum covering an energy range of approximately 1–17 keV in 48 pulse height analyzer (PHA) channels every 10 ms. In this work, we make no use of the lowest (0) and the highest (47) PHA channels; the PHA channels 1–46 (1–16 keV) are referred to as the total USA range. The spectral characteristics of the time series were studied by dividing the USA data into two energy bands, the PHA channels 1–10 (soft band) and 11–46 (hard band), corresponding to approximately 1–3 keV and 3–16 keV, respectively. A hardness ratio (HR), shown in the middle panel of Figure 1, is the ratio of the counting rate in the hard band over the soft band. The dates are given in MJD – 51,000.

2.2. RXTE/PCA Data

Unpublished PCA data of 1ES 1959+65 were obtained from the RXTE data archive. The RXTE/PCA observed 1ES 1959+65 12 times from 2000 July 28 through August 2,
and 5 times between 2000 September 1 and 6, with exposures of \(\sim 900\) s in the July-August observations and a few kiloseconds each in September. The STANDARD2 data were extracted using the HEASARC FTOOLS and filtered using the RXTE Guest Observatory Facility-recommended criteria (layer 1 only; for better signal to noise, PHA channels 0–27 [light curves only], or approximately 1–10 keV; Earth elevation angle greater than 10\(^\circ\); pointing offset less than 0.0; time since the peak of the last South Atlantic Anomaly passage greater than 30 minutes; and electron contamination less than 0.1). The light-curve and spectral data are from unit 2 (PCU 2) only since a faint background model is not yet available for unit 0 during gain epoch 5. The background models of epoch 4 were used. Light curves were extracted using the ftools SETSCREW (through the REX script) and LCURVE. The variable PCA background was modeled with PCABACKEST, which uses observations of X-ray–blank, high-latitude areas of the sky (Jahoda et al. 1996). Spectral fits were done using XSPEC v. 11.0.1 and response matrices generated by the ftool PCARMF.

The source was not detected with a better significance than \(\sim 2\alpha\) with the RXTE High-Energy X-Ray Timing Experiment (HEXTE) instruments. A likely explanation is the steep spectrum, derived from the PCA data, that falls below the HEXTE sensitivity.

2.3. Archival Data

An observation with the BeppoSAX instrument in 1997 (Beckmann 2000; Beckmann et al. 2002) resulted in a measured flux of \(1.3 \times 10^{-11}\) ergs cm\(^{-2}\) s\(^{-1}\) in the 2–10 keV band and a spectral index of \(\alpha = 1.64\). This flux is an order of magnitude fainter than the brightest flux measured here by the PCA. However, the results published by Beckmann (2000) should be treated with caution, as there is an apparent error in the value of the Galactic column adopted by them (they adopt \(10^{20}\) rather than \(1.027 \times 10^{21}\) atoms cm\(^{-2}\) adopted by us on the basis of the COLDEN program available as a part of the Chandra Proposal Planning Toolkit, based on relatively reliable 21 cm data). The analysis of the Position Sensitive Proportional Counter All-Sky Survey data by those authors implies that assuming a simple power law, the fitted absorption is \(1.6 \times 10^{21}\) atoms cm\(^{-2}\), somewhat larger than the Galactic value. However, this might be because the intrinsic X-ray spectrum steepens somewhat toward higher energies, as is often the case for other XBL-type blazars, and an assumption that the observed spectrum is a simple absorbed power law overestimates the fitted absorption. Note that the higher column density makes no difference in the results from the USA or RXTE data. The Einstein Slew Survey Sample of BL Lac Objects (Perlman et al. 1996) quotes a flux of \(3.65\) \(\mu\)Jy at 2 keV, which is \(\sim 40\%\) brighter than the BeppoSAX measurement.

We also extracted the ROSAT HRI archival data for this source and used the ROSAT All-Sky data. ROSAT HRI observed it on 1996 April 1; the observation lasted for a total of half a day, yielding about 2800 s of good data. The data were reduced in a standard manner, revealing that the net source counting rate was \(\sim 1.57\) counts s\(^{-1}\), with no indication of variability, but this is not too surprising given the short observation length. The conversion of the HRI count rate to flux is dependent on the source spectrum, which has to be assumed as there is essentially no spectral information in the HRI data. Since we do not know the soft X-ray spectrum, the fitted absorption is 1\(^{\pm}\)0.5\(\times 10^{21}\) cm\(^{-2}\). To obtain the conversion from the HRI counting rate to the observed flux, we used the PIMMS tool provided by HEASARC (and checked the results using XSPEC with the HRI effective area curve). Assuming the Galactic column of \(10^{21}\) cm\(^{-2}\), we obtain the 0.1–2.4 keV flux of \(5.5 \times 10^{-11}\) ergs cm\(^{-2}\) s\(^{-1}\) and 1–2 keV flux of \(2.2 \times 10^{-11}\) ergs cm\(^{-2}\) s\(^{-1}\). Assuming the fitted ROSAT value of \(1.6 \times 10^{21}\) cm\(^{-2}\), we infer the 0.1–2.4 keV flux of \(5.4 \times 10^{-11}\) ergs cm\(^{-2}\) s\(^{-1}\) and 1–2 keV flux of \(2.5 \times 10^{-11}\) ergs cm\(^{-2}\) s\(^{-1}\). In any case, this corresponds to (roughly) 2 mcram, which is lower than the 13 mcram level from the ROSAT All-Sky Survey Bright Source Catalogue (1RXS-B). This simply implies that during the ROSAT survey, 1ES 1959+65 was significantly brighter than during the Einstein, ROSAT HRI, or BeppoSAX observations, and the episodes of high flux as seen by the USA or RXTE observations described hereafter are not unique.

3. OBSERVATIONAL RESULTS

3.1. Flares in the USA and RXTE Data

The USA and RXTE observations of 1ES 1959+65 conducted from 2000 July through November show that the source was bright and variable in the X-ray band, with the X-ray spectrum significantly harder than observed during the periods of lower brightness. Specifically, these data show that during the last quarter of 2000, 1ES 1959+65 underwent an X-ray flare reaching the 13 mcram level in the 1–16 keV band on November 14 (MJD 51,863). Variability of a factor of \(\sim 6\) was detected within 20 days and a factor of \(\sim 3\) within 7 days (Fig. 1). By comparison, the peak flux detected by USA on Mrk 421 in 2000 reached approximately 40 mcram at maximum. Visual inspection of the USA (as well as the RXTE) light curves, and in particular of the largest flare, indicates that the source does not appear to vary significantly on timescales shorter than a day; thus, the variability is not undersampled. These observations show that 1ES 1959+65 was in a variable state for at least 4 months.

To complete the coverage of this flare, 3 data points from the RXTE all-sky monitor (ASM) were normalized to the Crab and added to the light curve in Figure 1; this shows the flare continuing to decrease. The varying spectral index and differences in the energy response of both instruments complicate the comparison. Nonetheless, the full-width at half-maximum region of the flare with these additional points spans 5 ± 1 days, and the doubling time is 2.5 days.

The PCA archival data obtained 2 months prior to the USA observations show 65% flux changes in 3.5 days; the highest observed value was \(f(2–10\text{ keV}) = 1.4 \times 10^{-10}\) ergs cm\(^{-2}\) s\(^{-1}\), but the peak value is unknown, since the observed maxima are at the endpoints of the observed period, when the source was falling or rising as shown in Figure 2. The USA light curve shows that the PCA did not cover the typical variation period, which appears to be greater than 4 days. The same figure shows a decreasing flux extending over 3.5 days, and after 30 days observations resume for 4 days where a steady increase of flux is seen. The flux did not change more than a few percent on timescales shorter than a day in the PCA data. In the three panels in Figure 3 it is
apparent that variations are larger in the harder bands in the decreasing part of the PCA observations.

3.2. Flux-Spectrum Correlations

The PCA data were used to perform spectral fits as a function of the flux in the 2–10 keV energy range. The data were fitted to a single power-law function with index $\alpha$, such that the photon flux $N(E) = N_0 E^{-\alpha}$ and the absorbing column $N_H = 10^{27}$ cm$^{-2}$.

The absorbed power-law model provides an adequate fit for all RXTE PCA pointings. Spectral indices were obtained for every observation, and in some cases, intraday observations where the estimated indices and fluxes were similar were added together to improve the significance on a daily timescale. The spectral fits are shown in Table 1 along with the month, day, and fraction of the day of the beginning of each observation.

The X-ray spectrum follows a "loop" in the spectral index-flux plane, as seen in Figure 4, and it is not surprising that the X-ray spectrum shows significant evolution during the flare given that there is a more rapid rise and drop in the hard X-ray band (see § 4 for discussion). The steepest spectrum in the PCA data was observed at the first observations in the declining phase July 28 ($\alpha = 1.42$), and the hardest spectrum was seen on two dates separated by a month and at a similar flux ($\alpha = 1.65$), which is a hint that the same physical mechanism is generating these flux variations. In the case of the USA observation, poorer photon statistics were added together to improve the significance on a daily timescale. The spectral fits are shown in Table 1 along with the month, day, and fraction of the day of the beginning of each observation.

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In the case of the USA observation, poorer photon statistics...
and an incomplete energy calibration limited the spectral study to an HR estimation plotted on the same Figure 1. The USA fluxes and HR are shown in Table 2. During the strongest flare, which started around \( \text{MJD} = 51,000 \), a 20% variation in the HR is observed.

### 4. DISCUSSION

#### 4.1. Doppler Boosting of the Flux of 1ES 1959+65

The electromagnetic emission in blazars is very likely to be Doppler-boosted (or beamed) toward the observer. In the radio regime, the evidence comes from superluminal expansions observed with VLBI. Similar superluminal expansions have now been seen in the optical band with the \( \text{HST} \) in nearby galaxies such as M87. Relativistic beaming is also required in order to avoid absorption of GeV photons by X-ray photons via the \( e^+e^- \) pair-production process. It is thus possible to use the X-ray variability data, as well as the fact that 1ES 1959+65 is an EGRET-detected BL Lac object, to establish a limit for the Doppler factor \( \delta \), with \( \delta \) defined in the standard way as \( \delta = \frac{1 - \beta \cos \theta}{1 + \beta \cos \theta} \), where \( \beta = v/c \), and \( \theta \) is the angle to the line of sight.

Assuming that the \( \gamma \)-rays and X-rays from 1ES 1959+65 are produced in the same region, it is possible to calculate the opacity for pair production \( \tau_{\gamma\gamma} \) from the source sizes inferred from the USA and \( \text{RXTE}/\text{PCA} \) data. The formula given by equation (3) in Mattox et al. (1993) for the optical depth for an outflow that is nonrelativistic in its comoving frame, as corrected by Madejski et al. (1996), is

\[
\tau = 2 \times 10^3 (1 + z)^2 \alpha (1 + z - \sqrt{1 + z})^2 \times h_{60}^2 T_5^{-1} \frac{F_{\text{keV}}}{\mu\text{Jy}} \left( \frac{E_x}{\text{GeV}} \right)^\alpha,
\]

where \( T_5 \) is the doubling time in units of \( 10^5 \) s and \( h_{60} \) the reduced Hubble constant in units of 60 km s\(^{-1}\) Mpc\(^{-1}\).

Using the parameters found in 1ES 1959+65 \( (z = 0.048, \alpha = 1.4, T_5 = 2.5, \text{and } F_{\text{keV}} = 40 \mu\text{Jy}) \), the opacity for 1 GeV photons would be \( \tau_{\gamma\gamma} \approx 20 \), and \( \delta > 1.6 \) is required in order to have \( \tau_{\gamma\gamma}(E_x > 1 \text{ GeV}) < 1 \). The limit on the required Doppler factor is less than that required for Mrk 421 \( (\delta = 5) \), which is not surprising since timescales of \( T = 0.5 \) days and a slightly higher luminosity are involved in that object.

Strictly speaking, the above \( \tau_{\gamma\gamma} \) argument for anisotropy only applies if the \( \gamma \)-ray−emitting zone is the same as the soft X-ray−emitting region; for now, we have no clear observational indications that this is the case. This is important since the jets are likely to be inhomogeneous. However, since \( \gamma \)-ray variability on timescales shorter than a day has been measured—and since there is an obvious correlation between X-rays and \( \gamma \)-rays that has been seen in objects such as Mrk 421 and Mrk 501—for such sources \( \delta > 1 \) can be deduced as well. Relativistic source motion, however, does not avoid the problem of the \( \gamma \)-ray pair production on the external, unbeamed photons likely to be present in the environments currently envisioned for the central engines of such sources.

The X-ray data presented above imply that the X-ray spectrum of 1ES 1959+65 hardens as the source brightens. This is often measured in BL Lac objects; a hardening of the spectrum when flares occur and a blueward shift of the peak of the synchrotron emission (and presumably higher energy inverse Compton emission) by factors that can be as large as 100 were measured in the cases of Mrk 501 (Pian et al. 1998), 1ES 1426+428, and PKS 0548−322 (Costamante et al. 2001). In the case of PKS 2005−489 (Perlman et al. 1999, hereafter PMSR99), a more moderate shift of a factor of 3 or less of the synchrotron emission was found.

#### 4.2. Synchrotron Models and Inferred Parameters

The spectral change is best illustrated as a correlation between flux and the photon index. Using PCA data this correlation is illustrated in Figure 4. Even though the two observations were separated by a month and are certainly related to two different flares, it is still interesting to compare this spectral evolution since the time series have similar rise and fall timescales, which are comparable to what is seen in the USA detector. It is thus likely that the two flares originate from a similar mechanism and that the correlation plot has some validity. The “clockwise motion” (shown with arrows) observed in the data for 1ES 1959+65 has also been seen in flares in Mrk 421 (Takahashi et al. 1996, 2000), in PKS 2155−304 (Sembay et al. 1993) and in H0323+022 (Kohmura et al. 1994), although in some cases counterclockwise patterns have also been seen (Mrk 501; Catanese & Sambruna 2000). The spectrum steepens more rapidly than the flux in the declining phase and hardens rap-
idly in the brightening phase, indicating that the variations of the hard X-rays occur faster than those in the soft X-rays both during the increase and the decrease of the brightness of the source. The spectral index change of 14% seen here is comparable to the 10% seen in Mrk 421. The variation observed in the flux-index plane can provide information about the acceleration process (Kirk & Mastichiadis 1999). Counter-clockwise patterns are expected when acceleration, variability, and cooling timescales are similar in a flare. In this case the acceleration process proceeds from low energy to high energy, changing the number of particles and making the softer energies vary first. Clockwise patterns, where the harder energies vary first, can be explained in flares where the variability and acceleration timescales are much less than the cooling timescale.

For a homogeneous emitting region, the radiative lifetime of a relativistic electron emitting synchrotron photons with energy $E_{\text{keV}}$ is (in the observer’s frame)

\[ \tau_{\text{sync}} = 1.2 \times 10^3 B^{-3/2} E^{-1/2} \delta^{1/2} \approx 1.2 \times 10^3 B^{-3/2} E^{-1/2} \delta^{1/2} \text{ s (Rybicki & Lightman 1979)}. \]

This should give some estimate of the magnetic field $B$, even though the extent to which the timescale of the flux decrease was due to the propagation of the signal through the source and what extent it was caused by the synchrotron cooling is not known. However, it is possible to measure the relative decrease of the flux $\Delta F/F$ in three energy bands in a time $\Delta T$ using the PCA data as seen in Figure 3. To estimate the timescale for a drop by a factor of 2 in each energy band, the measured timescale is divided by a factor $2\Delta F/F$. The factor $\Delta F/F$ is smallest for the lowest energy band (24%), as expected. It is now possible to find $\tau_{\text{sync}}(E)$ and $\tau_{\text{sync}}(12 \text{ keV})$ in the inferred values of $\tau_{\text{sync}}$. This means that the difference in the harder energies vary first, can be explained in flares where the variability and acceleration timescales are much less than the cooling timescale.

4.3. Continuous Emission: Knot Radiation?

The USA light curve exhibits a nonzero X-ray flux outside of the flaring events of a few $10^{-11}$ ergs cm$^{-2}$ s$^{-1}$ (or a few mcrab, taking 1 crab $\approx 1.7 \times 10^{-8}$ ergs cm$^{-2}$ s$^{-1}$ in the 2–10 keV band). The existence of a steady underlying emission in at least one other BL Lac object, Mrk 421, has been invoked in order to obtain meaningful fits with an exponential decay to X-ray flares (Fossati et al. 2000). It is intriguing to investigate if such steady flux could originate in more extended jets, such as those recently resolved by the Chandra and XMM telescopes, but aligned more closely to our line of sight and thus brighter. Such knots (or hot spots) in large-scale (hundreds of parsecs or more) jets have indeed been seen from radio to X-ray energies in many nonaligned sources, i.e., sources where the jet is sufficiently misaligned to allow us to resolve the structure of the jet; of course, such structures must also originate on a relatively large spatial scale as compared to the subparsec jets responsible for the rapid, day-scale variability. These knots are persistent structures visible on timescales of years, and individual spectral energy distributions have been established for knots in multiple sources such as M87 (Marshall et al. 2002) or PKS 0637−752 (Chartas et al. 2000). Their fluxes are usually a fraction of the flux arising from the unresolved core, but a closer alignment to our line of sight than for those sources resolved by Chandra and XMM would result in greater Doppler boost and could provide a continuous background seen in BL Lac–type objects, presumably the objects most closely aligned to our line of sight.

To verify if a flux of a few $10^{-11}$ ergs cm$^{-2}$ s$^{-1}$ could originate from aligned X-ray knots, some fluxes are taken from the literature for the cases where the angles to the line of sight ($\theta$) and Lorentz factors are available from radio observations. With this, it is possible to estimate the flux enhancement for $\theta \approx 0^\circ$ and the luminosity at a distance similar to that of 1ES 1959+65. The observed flux would depend on three factors:

1. A change in Doppler boost that multiplies the flux $F$ by $\delta/\delta^4$ (see Urry & Shaffer 1984 for luminosity conversions), where $\delta$ is the Doppler factor for the same source but with an angle $\theta = 0^\circ$. Thus, the amplification would be $(1 - \beta \cos \theta)/1 - \beta^4$. Knowing the Lorentz factor $\Gamma$, we can write $\delta^2 = (\Gamma^2 - 1)/\Gamma^2$.
2. The distance difference changes the flux by a factor of $(z/0.048)^2$.
3. A K-correction has to be applied since the bandwidth is narrowed, changing the flux by a factor of \((1+z)/1.048)^{-1}\), where \(\alpha\) is the spectral index. This effect is minimal compared to the previous ones, since in the X-ray band usually \(\alpha \geq 1\) for BL Lac spectra.

The final ratio \(A\) of the observed fluxes is then the product of these three terms.

4.3.1. Comparison with M87

The most popular scenario for the “parent population” of XBL-type blazars such as 1ES 1959+65 is that these objects are FR I type radio galaxies with their jets aligned close to our line of sight (see, e.g., Urry, Padovani, & Stickel 1991, hereafter UPS91). In contrast, PKS 0637−752 is a member of the intrinsically more powerful FR II class of objects. We thus believe that it is appropriate to compare 1ES 1959+65 with the FR I radio galaxy M87 (\(z = 0.004\)), a nearby object that has knots resolved in the X-ray band (Marshall et al. 2002) as well as in the radio and optical (e.g., Perlman et al. 2001 and references therein). In the most popular models, the X-ray radiation is either synchrotron radiation or inverse Compton (IC) scattering (see, e.g., Wilson & Yang 2002). From Biretta, Sparks, & Macchetto (1999, hereafter BSM99) we find that a relativistic jet model requires a bulk flow with Lorentz factor \(\Gamma \geq 6\) and a jet orientation within \(\theta \leq 19^\circ\) from the line of sight. We will use two sets of \((\theta, \Gamma)\) presented in BSM99 based on this model.

The set \((\theta, \Gamma) \sim (18^\circ, 6)\) is a possible configuration of the jet. These values may offer an explanation for the apparent lack of superluminal motion in M87 on parsec scales and were assumed in the model of Bicknell & Begelman (1996). From Marshall et al. (2002), the flux density at 1 keV summed over all the knots is approximately 380 nJy and \(\alpha = 1.46\). Using a distance of 16 Mpc for M87, \(H_0 = 60\) km s\(^{-1}\) Mpc\(^{-1}\), and \(q_0 = 0.1\), then \(A \sim 250\) would yield \(\sim 20 \times 10^{-11}\) ergs cm\(^{-2}\) s\(^{-1}\) (or 8 mcrab) of flux coming from the knots. In this case, using an angle of \(\sim 4^\circ\) would reduce the flux ratio by a factor of 8 (from 250), which would still keep the boosted knot radiation in the detectable range. An angle of \(4^\circ\) was used as well in this example since this is the mean value of angle to the line of sight used recently by Costamante & Ghisellini (2002) as an input parameter for the successful synchrotron self-Compton model those authors used to predict TeV candidate BL Lac objects.

The set of values \((\theta, \Gamma) \sim (16^\circ, 8)\) also presented in BSM99 and consistent with their results yield a flux coming from the knots that is 40 times smaller (or 0.2 mcrab). This configuration is thus less likely to result in a flux of the level that is being looked for.

4.3.2. Comparison with 3C 66B

3C 66B is a low-luminosity FR I radio galaxy (\(z = 0.0215\)). Its jet has knots resolved in the radio, optical, and X-ray band (Hardcastle, Birkinshaw, & Worrall 2001 and references therein). From Giovannini et al. (2001) we find that \(\theta\) seems well constrained (about 45\(^\circ\)) but that \(\Gamma\) is not (between 1 and 7).

The set \((\theta, \Gamma) \sim (45^\circ, 7)\) is a possible configuration of the jet that would come within the right order of magnitude to explain the origin of the X-ray radiation by IC scattering of seed photons from a hidden BL Lac object in the nucleus of 3C 66B. From Hardcastle et al. (2001), the flux density at 1 keV for the two brightest knots out of five (A and B) is approximately 10 nJy. A value of \(\alpha = 1.31\) is used, the one found for the jet, although the exact value does not matter much given the comparable redshift of 3C 66B to the one of 1ES 1959+65. The inferred value of \(A \sim 1.6 \times 10^5\) yields \(\sim 4 \times 10^{-9}\) ergs cm\(^{-2}\) s\(^{-1}\) (or 230 mcrab) of flux coming from the knots in the case of an angle at 0\(^\circ\). Using an angle of \(\sim 4^\circ\) would reduce the flux ratio by a factor of 2 only and make it still a bright steady source. The level of expected radiation in this case is actually so high that a similar object pointing closer to the line of sight would not be unnoticed in radio emission; thus, the \((\theta, \Gamma) \sim (45^\circ, 7)\) set seems quite extreme.

Taking a slower jet (\(\beta \sim 0.75\)) also mentioned in Hardcastle et al. (2001), the flux coming from the knots of 3C 66B would be a factor of 6 \(\times 10^4\) lower than in the previous case, significantly below the level of a few \(10^{-11}\) ergs cm\(^{-2}\) s\(^{-1}\).

The results from M87 and 3C 66B show that it is possible that radiation from knots moving close to the line of sight can generate a significant fraction of the baseline level of X-ray flux invoked in at least two cases, Mrk 421 and 1ES 1959+65, assuming they contain knots similar to those observed in M87 or 3C 66B. Recent VLBI observations show that the line-of-sight angle with Mrk 421 is constrained to be in the 0\(^\circ\)–30\(^\circ\) range (Giovannini et al. 2001); thus, a small angle to the line of sight is a possibility for this source. In a more general way, UPS91 predict that FR I radio galaxies should have jets with bulk flow speeds in the range \(\Gamma \sim 5–35\), with most near \(\Gamma \sim 7\), and they derive a critical angle \(\theta_{\text{crit}} \sim 10^\circ\) for the FR I/BL Lac division. If bright knots are common in BL Lac objects, then the amplification factors associated with the large Lorentz factors invoked here could generate a continuous background at the level mentioned above (assuming \(\theta\) is on the order of a few degrees only). A more thorough search for evidence of continuous emission in blazars and other signatures of boosted knot radiation is currently under way.

5. CONCLUSIONS

This paper reports X-ray data obtained with the USA and RXTE missions, as well as archival observations for the BL Lac object 1ES 1959+65. Variability on the timescale of a few days with a three-fold flux increase in the 1–16 keV band was observed with the USA detector; this is consistent with timescales that have been seen in the optical band at other times. The X-ray data also show a clear correlation of the X-ray flux with the spectrum, which becomes harder when the source is brighter. The data presented in this paper represent an important contribution to the study of this object, a potential TeV emitter, as long-term monitoring observations of it in the X-ray band were seldom performed in the past.

From the X-ray variability timescales, we estimate the Doppler factor of the jet to be greater than 1.6. The spectral variability data allow an estimate of the magnetic field that is lower (~milliGauss), and Lorentz factor of radiating electrons that is higher (~10\(^3\)), than those derived for the prototype TeV blazar Mrk 421. These parameters are more in line with those inferred for another possible TeV-emitting blazar, PKS 2005−489. The data also show the need for X-ray observations of 1ES 1959+65 longer than a few days in order to more accurately characterize the timescale of variability and its relationship with the spectrum. Also, since the source has been seen to be extremely variable in the optical
band, future multiwavelength observations of this object should include simultaneous optical observations.

We attempted to give a plausible scenario explaining the apparent baseline (“quiescent”) emission level seen here in 1ES 1959+65 but also possibly previously observed in Mrk 421. Jets containing bright knots, such as those seen by Chandra, but aligned more closely, within a few degrees of the line of sight (and therefore not resolvable as easily as jets with larger angles), and with Lorentz factors of ~10 have the required boosted flux to account for the “quiescent” X-ray emission in XBL-type BL Lac objects. Such emission would persist on timescales much longer than the duration of the flares that are much more likely to occur closer to the central object. This scenario can be confirmed by a more systematic study of knot X-ray luminosities in FR I objects, as well as the currently ongoing population study of the steady component of X-ray emission in BL Lac objects.

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