Measuring the mass accretion rates of Herbig Ae/Be stars with X-shooter *

M. A. Pogodin1,2,*, S. Hubrig3, R. V. Yudin1,2, M. Schöller4, J. F. González5, and B. Stelzer6

1 Central Astronomical Observatory at Pulkovo, St. Petersburg 196140, Russia
2 Isaac Newton Institute of Chile, Saint-Petersburg Branch, Russia
3 Leibniz-Institut für Astrophysik Potsdam, An der Sternwarte 16, 14482 Potsdam, Germany
4 European Southern Observatory, Karl-Schwarzschild-Str. 2, 85748 Garching, Germany
5 Instituto de Ciencias Astronomicas, de la Tierra, y del Espacio (ICATE), 5400, San Juan, Argentina
6 INAF-Osservatorio Astronomico di Palermo, Piazza del Parlamento 1, 90134 Palermo, Italy

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1 Introduction

Herbig Ae/Be stars (HAeBes) are pre-main sequence (PMS) objects of intermediate mass approximately from 2 to 10 $M_\odot$ (Herbig 1960; Finkenzeller & Mundt 1984). They range from spectral class F2–O9 and are surrounded by relict dust/gas accretion disks, reflecting themselves in the form of an IR excess as a result of thermal emission of circumstellar (CS) dust. For a number of stars, the disks can also be directly detected on images (e.g. Grady et al. 2001). CS environment of HAeBes possesses a complex structure. Gaseous streams are accreted from the equatorial disk onto the star and a stellar wind zone exists at higher latitudes. On the other hand, the disk wind, the so-called magnetocentrifugal wind, carries the excess of angular momentum away (e.g. Garcia et al. 2001). CS contribution to the stellar atmospheric spectrum provides information on physical processes in the envelope and on its interaction with the star. The character of this interaction is reasonably well studied for PMS objects with masses of about one solar mass and less (T Tauri-type stars or TTS). These objects have significant (several kG) dipole-type magnetic fields. The star/CS interaction in TTS is described in details. Our study shows that the majority of the calibrational relations can be applied to Herbig Ae/Be stars, but several of them need to be re-calibrated on the basis of new spectral data for a larger number of Herbig Ae/Be stars.

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** Corresponding author: e-mail: pogodin@gao.spb.ru

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but below the current sensitivity limits. The inner boundary of their accretion disks are likely to be located close to the stellar surface.

The existence of magnetically controlled disk accretion in HAes going through polar funnels was confirmed by Muzerolle et al. (2004), Mottram et al. (2007), and Grady et al. (2010). In any case, the magnetospheric radii of HAes with B \approx 100 \text{G} must be much smaller in comparison with those of TTS, where magnetic fields are of the order of several kG. As a result, we can expect the character of the star/CS interaction in HAes to be different from that of TTS. In our study aim at investigating the general picture of the star/CS interaction in HAes and at comparing it with that in TTS. One of the initial goals was also to study the variations of the accretion mass rates over the magnetic period. However, due to weather or technical constraints the observations were carried out mostly for targets with single magnetic field measurements, i.e. for targets with unknown magnetic phase curves. Clearly, an analysis of the temporal behaviour of different mass accretion rate indicators over the magnetic/rotation periods would provide fundamental clues on how magnetospheric accretion works. The targets for which magnetic/rotation periods have been determined in our previous studies include HD 97048 and HD 101412 with only one X-shooter observations, respectively, HD 150193 with two observations, and HD 176386 with four observations (Hubrig et al. 2011b, 2011a).

One of the most important parameters characterizing the properties of the star/CS interaction is the mass accretion rate $\dot{M}_{\text{acc}}$. It determines strongly the disk structure and dynamics (e.g., D’Alessio et al. 1999, 2001). Measurements of $\dot{M}_{\text{acc}}$ are well-developed for TTS, based on a number of empirical spectral diagnostics (e.g., Oudmaijer et al. 2011, Rigliaco et al. 2011 and references therein).

Three main purposes of our work presented here are:

1. To determine $\dot{M}_{\text{acc}}$ in a sample of HAes whose magnetic fields have been detected (Hubrig et al. 2009, 2011b) using calibrations previously obtained for TTS and brown dwarfs (BDs). Magnetic HAes appear to be the most suitable targets for such a study because their disk accretion process is expected to better fit the MA scenario in comparison with non-magnetic stars.

2. To test the applicability of different calibrations for measuring $\dot{M}_{\text{acc}}$ of HAes and to select the best spectroscopic tracers of accretion necessary to provide accurate accretion rates for systematic studies of large samples of HAes.

3. To search for variability of $\dot{M}_{\text{acc}}$ in HAes with more than one available spectrum. This variability can be twofold: a) a real change in the accretion process, and b) a modulation of $\dot{M}_{\text{acc}}$ by the rotating magnetosphere and a cyclic screening of the stellar limb by local funnel streams.

## 2 Mass accretion rate calibrations

### 2.1 Calibrations for TTS and BDs

At present, a number of empiric spectral indicators are available for the determination of $\dot{M}_{\text{acc}}$ for TTS and BDs.

#### 2.1.1 Model method

The basic accretion indicator is the additional emission in the continuum superimposed onto the stellar spectrum and calculated on the basis of MA models for these objects (e.g., Calvet & Gullbring 1998, Gullbring et al. 2000, Ardila et al. 2002). It is assumed that the overwhelming bulk of this emission originates in funnel flows inside the magnetosphere and in the shock zone on the stellar surface in a compact region. This assumption is rather plausible since the emissivity in the continuum is proportional to $N_e^2$ (where $N_e$ is the electron density), and the magnetospheric streams consist of the densest CS matter. The gas density in parts of the envelope (disk and wind) drops very quickly with distance from the star, and its contribution to the additional continuum emission is negligible. In the observed spectra, this emission is seen as veiling of the atmospheric lines and as an excess radiation at wavelengths below the Balmer jump. The theoretical models allow to construct a quantitative relation between $\dot{M}_{\text{acc}}$ and the veiling plus UV excess. Yet, veiling and UV excess frequently prove not to be convenient for these measurements, due to the large extinction in the regions where the targets are located.

#### 2.1.2 L - type calibrations

Alternatively, empirical calibrations have been determined between emission lines and mass accretion.

All emission lines in the spectrum of a PMS object are of multi-component origin. They are formed not only in the magnetospheric streams, but also in the disk and in the wind. The source function determining the emissivity in a line drops with the distance from the star much slower than the emissivity in the continuum (for example, see Pogodin 1986, 1989 for Balmer emission lines). However, according to the MA models, such parameters as the mass of the star, mass of the disk, $\dot{M}_{\text{acc}}$ and mass loss rate $\dot{M}_{\text{wind}}$ are likely to be interdependent. It has been established that the accretion luminosity ($L_{\text{acc}}$) is clearly correlated with luminosity in some emission lines ($L_{\text{line}}$). A number of empirical calibrations have been published so far. The dependencies are approximated by expressions of the type:

$$\log(L_{\text{acc}}) = a \log(L_{\text{line}}) + b,$$

$$\dot{M}_{\text{acc}} = \frac{L_{\text{acc}}R_*}{GM_*},$$

where $M_*$ and $R_*$ are the stellar mass and radius, respectively, and $a$ and $b$ are coefficients determined separately for each spectral line (e.g., Fang et al. 2009, Dahm 2008, Herczeg & Hillenbrand 2008, Gatti et al. 2008, Natta...
et al. 2004, Mohanty et al. 2005. The calibrations are constructed on the basis of spectroscopic observations of the chosen line in a representative sample of targets, whose $\dot{M}_{\text{acc}}$ values were determined beforehand using the model method.

### 2.1.3 F-type calibrations

Besides of that, another type of empirical calibrations exists that is based on a direct correlation between $\dot{M}_{\text{acc}}$ and the flux in the emission line ($F_{\text{line}}$). The dependencies are expressed as:

$$\log(\dot{M}_{\text{acc}}) = c \log(F_{\text{line}}) + d,$$

where $c$ and $d$ are coefficients (e.g. Herczeg & Hillenbrand 2008). The F-type calibrations are constructed analogous to the case of L-type calibrations where the values of $\dot{M}_{\text{acc}}$ are derived beforehand using the model method or from other types of calibration.

### 2.1.4 Hα10% - indicator

There is also a separate accretion indicator, the Hα10%. It is defined as the full width W of the emission Hα profile at the level of 10% of the maximum intensity (White & Basri 2003). Natta et al. (2004) extended this indicator to very low-mass objects, where $\dot{M}_{\text{acc}}$ was determined by fitting the observed Hα profiles with predictions of MA models (Muzerolle et al. 2001). The calibration is expressed in the form:

$$\log(\dot{M}_{\text{acc}}) = c_1 \log(W(\text{H}\alpha10\%)) + d_1,$$

where $c_1$ and $d_1$ are coefficients and W(Hα10%) is given in (km/s) (Natta et al. 2004).

### 2.2 Calibrations for HAeBes

At present, calibrations constructed specifically for HAeBes are small in number.

Donehew & Brittain (2011) derived $\dot{M}_{\text{acc}}$ for a considerable sample of HAeBes using the model method, where the modeling was applied to the emission Balmer jump $\Delta D_B$. They calculated $\dot{M}_{\text{acc}}$ in dependence of the emission Balmer jump for more than 30 HAeBes using the models and methods developed by Calvet et al. (1998), Gullbring et al. (2000), and Muzerolle et al. (2004). The model of Muzerolle et al. (2004) takes into account the main physical processes predicted by the MA model such as: a ballistic infall of material from the disk onto the stellar surface along accretion columns, heating of the photosphere due to its shock interaction with the accretion streams and releasing soft X-rays that are absorbed by the surrounding material. This material then emits optical and UV radiation as it thermalizes. In the model, the flux from accretion is calculated for three different regions: the shock region, the heated photosphere, and the accretion columns, to get the overall flux from accretion, which is a function of $M_\ast$, $R_\ast$, and $\dot{M}_{\text{acc}}$. The resulting spectral energy distribution (SED) including both the accretion and the stellar flux is used to calculate the expected emission Balmer jump $\Delta D_B$ for different $\dot{M}_{\text{acc}}$ and stars of different spectral types. The value of $\Delta D_B$ was defined as the difference in $D_B$ (Balmer jump) for a given star from that of a standard (non-accreting) star of the same spectral type, where $\Delta D_B$ is the difference in the magnitudes at both sides of the discontinuity. The relationship between $\Delta D_B$ and $L_{\text{acc}}$ calculated by Muzerolle et al. (2004) and between $L_{\text{acc}}$ and $\dot{M}_{\text{acc}}$ given in Sect. 2.1.2 were used by Donehew & Brittain (2011) to determine $\dot{M}_{\text{acc}}$ of their program stars with different $M_\ast$ and $R_\ast$. They used the expression:

$$\Delta D_B = 2.5 \log \frac{F_\star + F_A}{F_\star}.$$

The observational $\Delta D_B$ for all targets have been measured in the observed spectra normalized to the flux $F_\chi$ at 4000 Å using the spectra of standard stars and the Castelli-Kurucz LTE models (Castelli & Kurucz 2003) of stellar atmospheres with corresponding $T_{\text{eff}}$ and log $g$. As it turned out, the observational $\Delta D_B$ demonstrated a very convincing correlation with theoretical $\dot{M}_{\text{acc}}$. This result is likely to be related to the fact that the $M_\ast/R_\ast$-ratio is practically the same in PMS stars from late-B to K types with a dispersion of ±0.1 dex. We have built the dependence between $\Delta D_B$ and $\dot{M}_{\text{acc}}$ using the data taken from Columns 2 and 3 of Table 3 in Donehew & Brittain (2011) and fitting them by a polynomial. This relation is shown in Fig. 1 of this paper and used in our study.

Donehew & Brittain (2011) measured also the equivalent widths of the Brγ emission line and constructed the empirical correlation between $\log(L_{\text{acc}})$ (determined from $\Delta D_B$) and $L_{\text{Brγ}}$. They concluded that this relationship is different for HAes and for earlier type HAeBes (HBes)
and that the correlation for HAeBs is in satisfactory agreement with the calculation constructed earlier by Calvet et al. (2004) for classical TTS (CTTS) and IMTTS. Therefore, it appears that the Brγ emission line is probably a reliable tracer of $\dot{M}_{\text{acc}}$ in HAeBs. One of the goals of our study is testing the reliability of additional available spectral indicators.

Also Garcia Lopez et al. (2006) determined $\dot{M}_{\text{acc}}$ for 36 HAeBs using a similar calibration for TTS. Eight objects are in common between their target list and the list of Donehew & Brittain (2011). For four objects the agreement in the derived $\dot{M}_{\text{acc}}$ values is rather good with $\Delta \log M_{\text{acc}}$ ranging from 0.01 to 0.23 dex. Four other objects demonstrate significant differences from 0.7 to 1.2 dex. As a whole, the difference between the results of these two works makes up $-0.18 \pm 0.75$ dex, where the error is the standard deviation. Taking into account that the accuracy of an individual estimate is of the order of $\pm 0.4-0.5$ dex in both papers, the conclusion can be made that the results obtained in these two works demonstrate no significant difference.

During the time when our work was in preparation one additional paper was published by Mendigutia et al. (2011a), where a similar model method was used to determine mass accretion rates in 38 HAeBes and the L-type calibration relationships were constructed for several emission lines. We compared the results of this work with those obtained in Donehew & Brittain (2011) and Garcia Lopez et al. (2006). The Mendigutia et al. (2011a) sample has overlap with the target list of Donehew & Brittain (2011) – 14 stars – and with that of Garcia Lopez et al. (2006) – 13 stars. We see that $\dot{M}_{\text{acc}}$ of about half the targets of Mendigutia et al. (2011a) are estimated to be systematically higher in comparison with those obtained in the two other studies (+0.74 ± 0.44 dex for Donehew & Brittain 2011 and +1.67 ± 0.83 dex for Garcia Lopez et al. 2006). Only upper limits of $M_{\text{acc}}$ have been obtained in Mendigutia et al. (2011a) for the second half of their targets. The differences in their $M_{\text{acc}}$ determinations with those obtained in the two other works account for more than $-0.74 \pm 0.56$ dex compared to the work of Donehew & Brittain (2011) and $-0.50 \pm 0.24$ dex compared to the work of Garcia Lopez et al. (2006).

Trying to understand the possible cause of these discrepancies, we compared the values of the emission Balmer jump $\Delta D_B$ measured in Mendigutia et al. (2011a) by a photometric method and in Donehew & Brittain (2011), where a spectroscopic method was used. We have obtained the difference $0.17 \pm 0.17$. Taking into account the different dates of observations and different methods of measurements and their errors ($\pm 0.07$ for Mendigutia et al. 2011a and $\pm 0.03$ for Donehew & Brittain 2011), we can conclude that the differences in these measurements are not essential.

Thus we assume that some systematic distinctions can be expected in a model calculating the relationship between $M_{\text{acc}}$ and $\Delta D_B$ carried out in the different studies. We note that $M_{\text{acc}}$ values determined in Mendigutia et al. (2011a) appear significantly overestimated: ten objects in their study present rates ranging from $10^{-4}$ to $10^{-6}$ $M_\odot$/yr. Such large mass accretion rates can be expected only in protostars or in such unique objects as FUors. In addition, according to the stellar masses and ages presented in Table 1 of Mendigutia et al. (2011a), six out of ten targets of the spectral type B8 – A2 are located on the HR diagram close to the main sequence where the accretion rates already become smaller than at earlier stages of the PMS evolution. Generally, the masses of accretion disks around HAeBes range from 0.01 to 0.1 $M_\odot$ (Hillenbrand et al. 1998, Henning et al. 1998, Natta et al. 2000). With mass accretion rates of the order of $10^{-4}$ to $10^{-6}$ $M_\odot$/yr, the accretion disk would be completely dissipated already after $10^2 - 10^3$ years. This time interval is much smaller than the time of the whole PMS evolution stage of objects with masses of $2 M_\odot$, which is about $10^6$ years according to Palla & Stuhler (1993).

Based on these results, the question arises whether calibrations determined for lower mass pre-main sequence stars, T Tauri stars, can be employed to estimate $M_{\text{acc}}$ in HAeBes. Clearly, applicability of such calibrations for $M_{\text{acc}}$ in HAeBes demands a special examination. It can be assumed that identical calibrations are applicable to HAeBes provided that the character of the disk/magnetosphere and the disk/wind interaction in TTS and HAeBes is similar and that the contribution of the CS regions to the whole CS spectrum formed in TTS and HAeBes is comparable too. As was discussed in Sect. II the size of the magnetosphere and the value of $B$ is expected to be much smaller in HAeBes in comparison with TTS. It is not clear yet whether these differences have an impact on the relations between $M_{\text{acc}}$ and the emerging radiation in spectral lines.

## 3 Observations and data analysis

The observations were performed in service mode with the X-shooter spectrograph mounted on the 8m UT2 of the VLT. X-shooter allows to obtain spectral data simultaneously over the entire wavelength range from the near-UV to the near-IR in three arms (the UVB-arm covering the range 300–590 nm, the VIS-arm 550–1000 nm, and the NIR-arm 1.0–2.5 μm). The observations were performed with the

| Target stars for which X-shooter data were obtained during our observing run. Spectral type and photometric data were taken from the SIMBAD database. |
| Object | Spectral type | V | R | J | H | $A_V$ |
|--------|---------------|---|---|---|---|------|
| HD 97048 | A0pshe | 8.46 | 8.50 | 7.27 | 6.67 | 1.00 |
| HD 100546 | B9Vne | 6.80 | 6.70 | 6.43 | 5.96 | 0.10 |
| HD 101412 | B9.5V | 9.29 | 9.30 | 8.64 | 8.22 | 0.54 |
| HD 135344B | F4–F8 | 7.85 | 7.83 | 7.31 | 6.67 | 0.10 |
| HD 150193 | A1Ve | 8.88 | 8.90 | 6.95 | 6.21 | 1.60 |
| HD 176386 | B9Ve | 7.30 | -- | 6.90 | 6.75 | 0.60 |
| HD 190073 | A2IVpe | 7.82 | 7.80 | 7.19 | 6.65 | 0.12 |
| PDS 2 | F2 | 10.73 | -- | 10.01 | 9.68 | 0.55 |
Table 2  Stellar parameters of the targets. The majority of the data are taken from Hubrig et al. (2009). Further data is from: a – Catala et al. (2007); b – Montesinos et al. (2009); c – Coulson & Walter (1995); d – this work.

| Object       | $T_{\text{eff}}$ | $\log g$ | $M_*/M_\odot$ | $R_*/R_\odot$ | $v \sin i < B_z$ | $< B_z >$ |
|--------------|------------------|----------|----------------|---------------|------------------|-----------|
| HD 97048     | 10000            | 4.0      | 2.5            | 2.2           | 140              | 188±47    |
| HD 100546    | 10500            | 4.5      | 2.5            | 1.5           | 65               | 89±26     |
| HD 101412    | 9500             | 4.0      | 2.5            | 2.7           | 5                | -454±42   |
| HD 135344B   | 6250$^c$         | 4.0$^c$  | 1.36           | 1.9           | 70               | -37±12    |
| HD 150193    | 9000             | 4.0      | 2.2            | 2.0           | 100              | -252±48   |
| HD 176386    | 10000            | 4.0      | 2.7            | 2.3           | -121±35          |           |
| HD 190073    | 9250             | 3.5      | 2.85$^a$       | 3.6$^a$       | 12               | 104±19    |
| HD 190073$^b$|                 |          |                |               |                  |           |
| PDS 2        | 7000$^d$         | 4.0$^d$  | 2.5            | 1.6           | 30$^d$           | 103±29    |

Table 3  Non-accreting stars used as flux standards

| Object       | Spectral type | MJD  | $V$  |
|--------------|---------------|------|------|
| HD 100604    | F2V           | 55352.008 | 7.69 |
| HD 100627    | F6IV/V        | 55408.332 | 8.49 |
| HD 100926    | A3III/IV      | 55466.170 | 9.75 |
| HD 100928    | A0IV          | 55648.147 | 9.52 |

highest possible spectral resolution, i.e. $R$ is $\sim$9100 in the
UVB-arm, $\sim$17 400 in the VIS-arm, and $\sim$11 000 in the
NIR-arm (D’Odorico et al. 2006). The data were reduced
using the X-shooter pipeline (Version 1.1.0) following the
standard steps. For more details see Modigliani et al. (2010).

Due to a very high efficiency of the X-shooter spectrograph
the signal-to-noise ratio (S/N) of 300–500 was achieved
during exposure times ranging from 13–15 s for the bright-
est targets to 450 s for the faintest target PDS 2.

26 spectra of eight HAes with magnetic field detections
were obtained during 13 nights distributed between May
and September 2010. After each science exposure, telluric
standards were automatically observed in Obseratory time
at S/N$\sim$100. They are usually main-sequence hot stars or
solar analogs. Further, the data package delivered by ESO
includes a number of spectrophotometric standards. The
flux standards used in our work are presented in Table 3.

The original observation request foresaw several observ-
ations per target over the rotation period, but due to tech-
nical problems with X-shooter and mediocre weather con-
tions, the program was only partially completed. The
targets were chosen from the sample of Herbig Ae/Be stars
investigated previously by (Hubrig et al. 2009). The list of
X-shooter targets is presented in Table 1 and their stellar
parameters together with the detected longitudinal magnetic
fields are summarized in Table 2. Since the star PDS 2 was
less intensively studied in the past than other targets, we
estimated atmospheric parameters of this object by compar-
between the observed spectrum in the range $\lambda \lambda 4460$–
4500 Å and the synthetic spectrum computed using the code
SYNTH+ROTATE (Piskunov 1992). The best model fit is
presented in Fig. 2 with parameters: $T_{\text{eff}} = 7000$ K, $\log g =
4.0$, and $v \sin i = 30$ km/s. The complete list of observing
dates is given in Table 5.

The data analysis concentrated on the emission lines and
the emission Balmer jump that can be considered as accre-
tion indicators. To remove telluric features, we used the tel-
laric standards observed immediately after each observation
at a similar zenith distance. The equivalent widths (EWs) of
emission lines were measured after subtraction of stellar at-
tmospheric profiles from the observed ones. The synthetic at-
tmospheric profiles were calculated using the computer code

Fig. 2  Comparison of the mean observed spectrum of
PDS 2 in the range of $\lambda \lambda 4460$–4500Å with the synthetic
spectrum corresponding to the best fit. Top panel: The syn-
thetic spectrum. Bottom panel: The overplotted observed
spectrum (solid line) and the synthetic spectrum (dotted
line)
SYNTH+ROTATE (Piskunov 1992) and the standard LTE Kurucz models for the corresponding values of the stellar parameters (see Table 2). The synthetic atmospheric line profiles were also smoothed in accordance with the spectral resolution and $v \sin i$. For the near-IR hydrogen lines (Brγ, Paβ, and Paγ), the excess of radiation was taken into account by comparing the photometric magnitudes of the target stars with those for unreddened stars of corresponding spectral type and a correction for $A_V$ using the standard extinction law. All necessary data for this step are presented in Table 1. The EWs of the Na I D lines were measured as a sum value for both components $D_1$ and $D_2$, where the water vapour absorption lines and the interstellar (IS) narrow components were removed prior to this step. The IS components of the Na I D lines were cleaned by simple cutting. This procedure leads to a small underestimation of the EW, but it does not account for more than 10% ($\sim 0.1 \text{ dex}$) in the value of $\log M_{\text{acc}}$. The emission lines of the IR Ca II triplet are blended by weak emission components of Pa13, Pa15, and Pa16. But their contribution to the EWs of the Ca II lines can be taken into account by polynomial interpolation of the EW(Pa-lines) using synthetic spectra based on the unblended Pa12, Pa14, and Pa17 lines (Fig. 6). We used standard LTE Castelli-Kurucz models (Castelli & Kurucz 2008) to determine fluxes at different wavelengths corresponding to lines chosen as accretion indicators for all program stars. The values of their stellar masses $M_*$ and radii $R_*$ were taken from Table 2 to calculate luminosities in the line-indicators.

The procedure of measuring the emission Balmer jump $\Delta D_B$ is illustrated in Fig. 3 using HD 100546 as an example. The ratio $r$ of the initial unreddened spectra of the target (HD 100546) and the standard star of spectral type A0 (HD 100928) is approximated polynomially between 4000 Å and 4600 Å. The polynomial is extrapolated to 3640 Å and the ratio of ordinates $r_2$ and $r_1$ allows to determine the value of $\Delta D_B = 2.5 \log (r_2/r_1)$. This procedure is identical to that used in the work of Donehew & Brittain (2011).

### 4 Spectral accretion indicators

We examined 13 spectral diagnostics of $M_{\text{acc}}$ that have previously been used in studies of TTS and BD. The basic indicator was the emission Balmer jump $\Delta D_B$, calculated by the method described in Sect. 3. We used the relation between $\Delta D_B$ and $M_{\text{acc}}$ presented in Fig. 1. Eight indicators, based on the measurements of $L_{\text{acc}}$, include He I 5876, Brγ, Na I D, Paβ, Paγ, and Ca II λ8542. We considered also three indicators where flux in an emission line was used for the direct determination of $M_{\text{acc}}$ according to eq. 3: He I λ5876, Ca II λ8542, and Ca II λ8662. Further in the text we use the abbreviations “(L)” and “(F)” to mark indicators based on the measurements of $L_{\text{acc}}$ and flux, respectively (e.g., Paγ (L), Ca II λ8662 (F), etc.). The last indicator is Hα10%, which was calibrated by Natta et al. (2004). All these diagnostics are presented in Table 4.

Selected regions of the observed spectra of the program stars are presented in Figs. 4 to 7.

#### 4.1 Hα emission

The emission Hα profiles vary strongly from object to object (Fig. 4). HD 150193 and HD 190073 demonstrate P Cyg-type profiles indicating the existence of a stellar wind between the star and the observer. The Hα profiles of HD 100546 and HD 101412 are double-peaked with a central absorption shifted relative to the stellar rest wavelength of the line. The Hα profile of HD 97048 is double-peaked too, but of a complex structure, and single emission profiles are observed in the spectra of HD 135433B and PDS 2.
Fig. 4  Normalized Hα profiles in the spectra of studied HAes. For HD 135344B, HD 150193, HD 176386, HD 190073, and PDS 2 the displayed profiles are observed on dates MJD55357.329-371, 55352.047, 55465.138, 55446.078, and 55410.343, respectively. Profiles of HD 135344B and PDS 2 are shown in comparison with synthetic profiles calculated with the computer code SYNTH+ROTATE of (Piskunov, 1992). The LTE models of Castelli & Kurucz (2003) are used for the stellar parameters listed in Table 2. In the case of PDS 2, the value $v \sin i = 30$ km/s was used to calculate the model profile.
Table 4  Mass accretion rate diagnostics that have been used in previous studies of TTs and BDs and are analyzed in this work

| Reference            | Diagnostic                      | Mass range \([M_*/M_\odot]\) |
|----------------------|---------------------------------|-------------------------------|
| Fang et al. (2009)   | Hα (L), Hβ (L), He I λ5876 (L) | 0.1 – 2.0                     |
| Herczeg & Hillenbrand (2008) | Na I D (L), Ca II λ8542 (F), Ca II λ8662 (F), He I λ5876 (F) | 0.1 – 1.0                      |
| Dahm (2008)          | Paβ (L), Paγ (L), and Ca II λ8542 (L) | 0.5 – 2.0                     |
| Natta et al. (2004)  | Paβ (L), Paγ (L)                | 0.01 – 0.1                    |
| Gatti et al. (2008)  | Paβ (L), Paγ (L)                | 0.1 – 1.0                     |

As can be seen in Fig. 4, the Hα emission is practically invisible in the spectrum of the B9e star HD 176386. This fact was already discussed by other authors. Bibo et al. (1992) characterized this object as a higher-mass analog of so-called weak-line T Tauri stars (WTTS) with already dispersed accretion disks. On the other hand, Grady et al. (1993) report that signatures of a matter infall onto the star is observed in the UV spectrum of HD 176386. If the disk of HD 176386 is already dispersed, calibrations derived for the magnetospheric accretion scenario are no longer valid because the balance of the emission from different parts of the circumstellar components (magnetosphere, disk, wind) changes. In any case, the standard model of magnetospheric accretion from the disk can hardly be applied to this object. Therefore, any empirical spectral calibration based on the MA model cannot be used.

4.2 Other emission lines

Fig. 5 illustrates the spectra of the targets in the region of the He I λ5876 and Na I D lines, and Fig. 4 in the region of the IR Ca II triplet. The emission profiles of He I λ5876 are in general single-peak or double-peak type in the spectra of the HAeBes, but in some cases a deep redshifted absorption overlaps the red emission wing (see for more details Bohm & Catala 1995, Mendigutia et al. 2011b). An example of such a feature is seen in the spectrum of HD 97048 (Fig. 5). This line originates in the highest temperature regions of the CS envelope involving the inner boundary of the accretion disk interacting with the magnetosphere, the matter streams infalling onto the star, and the region of the impact of the streams on the stellar surface. These regions are rather compact in size and screening of the stellar limb by the accreted flows can be an important factor in forming the observed line profile. This screening leads to the appearance of a redshifted absorption component of the profile. Therefore, we have to be careful in using the He I emission for the \(\dot{M}_{\text{acc}}\) determination. At the time when the infalling stream screens the stellar limb, the \(\dot{M}_{\text{acc}}\) measurement can be considerably underestimated.

The types of the Na I D line profiles (Fig. 5) as well as of the IR Ca II triplet profiles (Fig. 5) in the spectra of the targets are very diverse. Even in objects of similar spectral types (A2–B9) are the intensities of these emission lines completely different. The emission in the Na I D lines is clearly visible and the EW can be measured in three objects HD 97048, HD 100546, HD 150193, and HD 190073 in the spectral region containing near-IR Ca triple lines. The positions of the Pa lines are indicated in the upper panel, and the positions of the Ca II lines in the lower panel. Strong emission near-IR Ca triplet lines appear only in two HAe stars, HD 150193 and HD 190073.
As for the Ca II triplet, the EW can be measured only in HD 190073 and HD 150193. It is remarkable that in our sample HD 190073 is the star with the strongest intensities of the Na I D and IR Ca II lines. This object demonstrates the clearest developed P Cyg-structure of the Hα profile (Fig. 4). The second object with an Hα profile of P Cyg-type is HD 150193. It shows also rather intense emission in Na I D and IR Ca II. This fact leads us to the hypothesis that the appearance of considerable emission in the Na I D and IR Ca II lines is related to the presence of a wind zone between the star and the observer. The signatures of the matter outflow are clearly seen in the Na I D lines of HD 190073 and HD 150193 in form of the P Cyg-structure.

4.3 Balmer jump

Fig. 7 illustrates the normalized spectra of the targets in the region of the Balmer jump in comparison with spectra of non-accreting stars of similar spectral types. Castelli-

(HD 190073, HD 150193, and HD 101412). As for the Ca II triplet, the EW can be measured only in HD 190073 and HD 150193. It is remarkable that in our sample HD 190073 is the star with the strongest intensities of the Na I D and IR Ca II lines. This object demonstrates the clearest developed P Cyg-structure of the Hα profile (Fig. 4). The second object with an Hα profile of P Cyg-type is HD 150193. It shows also rather intense emission in Na I D and IR Ca II. This fact leads us to the hypothesis that the appearance of considerable emission in the Na I D and IR Ca II lines is related to the presence of a wind zone between the star and the observer. The signatures of the matter outflow are clearly seen in the Na I D lines of HD 190073 and HD 150193 in form of the P Cyg-structure.

4.3 Balmer jump

Fig. 7 illustrates the normalized spectra of the targets in the region of the Balmer jump in comparison with spectra of non-accreting stars of similar spectral types. Castelli-
Fig. 7 Normalized spectra of our targets in the region near 3600–4400 Å. The normalization was made to the flux $F(\lambda)$ at the level of the model continuum at $\lambda = 4000$ Å. Spectra of standards of similar spectral type are given for comparison (dashed lines).

Kurucz LTE (Castelli & Kurucz, 2008) models were applied for normalization of the observed spectra of non-accreting stars to flux $F_{\lambda}$ at 4000 Å at the continuum level. The $r$-dependencies (see Sect. 3 and Fig. 3) were used to construct the normalized spectra of the targets.

The EWs of all studied lines, the Hα10% width and $\Delta D_B$ are listed in Table 5 together with the observing dates and magnetic/rotation phases calculated according to the ephemerides presented in Hubrig et al. (2011a, 2011b). Mean accuracies of the measurements are presented in Table 6. The errors of the individual estimates of the equivalent widths were computed using a program that takes into account the width and the maximum intensity of an emission line as well as the S/N ratio of the spectrum near the mea-
Table 5  Observing dates, magnetic/rotation phases calculated according to the ephemerides presented in Hubrig et al. (2011a, 2011b), and measurement results of EWs, \( W(H\alpha10\%) \), and \( \Delta D_B \) for HAes in our sample.

| Object      | MJD     | Magn. phase | Br\( \gamma \) | Pa\( \beta \) | Pa\( \gamma \) | EW [\( \text{Å} \)] Ca II 8662 | Ca II 8542 | H\( \alpha \) | Na I D | He I 5876 | H\( \beta \) | W(\( H\alpha10\%) \) | \( \Delta D_B \) |
|-------------|---------|-------------|----------------|-------------|--------------|-------------------------|-------------|-----------|--------|-----------|---------|----------------|-------------|
| HD 97048    | 55352.000 | 0.43         | 10.30          | 21.90       | 9.12         | – –                     | – –         | 30.8      | –       | 0.0525    | 1.58     | 604            | 0.25        |
| HD 100546   | 55352.022 | 0.36         | 2.27           | 3.88        | 1.47         | – –                     | – –         | 10.3      | –       | 0.43      | 1.57     | 496            | 0.15        |
| HD 101412   | 55351.992 | 0.71         | 5.79           | 7.77        | 3.63         | – –                     | – –         | 14.5      | –       | 0.32      | 2.16     | 501            | 0.15        |
| HD 135344B  | 55352.034 | 0.22         | 3.21           | 4.10        | 1.72         | – –                     | – –         | 11.1      | –       | 0.26      | 2.00     | 501            | 0.15        |
| HD 150193   | 55352.047 | 0.59         | 8.83           | 7.66        | 4.93         | 0.87                     | 1.61        | 11.9      | 0.253    | 0.035     | 1.69     | 560            | 0.25        |
| HD 176386   | 55329.362 | 0.27         | 2.27           | 1.88        | 0.47         | – –                     | – –         | 13.6      | 0.340    | 0.145     | 1.69     | 560            | 0.25        |
| HD 190073   | 55329.397 | 0.22         | 3.21           | 4.10        | 1.72         | – –                     | – –         | 11.1      | –       | 0.26      | 2.00     | 501            | 0.15        |
| PDS 2       | 55375.418 | 0.16         | 0.627          | 0.358       | – –          | 3.16                     | – –         | 0.015     | 0.015    | 0.001     | 0.015    | 0.02           | 0.03        |
| HD 97048    | 55352.000 | 0.43         | 10.30          | 21.90       | 9.12         | – –                     | – –         | 30.8      | –       | 0.0525    | 1.58     | 604            | 0.25        |
| HD 100546   | 55352.022 | 0.36         | 2.27           | 3.88        | 1.47         | – –                     | – –         | 10.3      | –       | 0.43      | 1.57     | 496            | 0.15        |
| HD 101412   | 55351.992 | 0.71         | 5.79           | 7.77        | 3.63         | – –                     | – –         | 14.5      | –       | 0.32      | 2.16     | 501            | 0.15        |
| HD 135344B  | 55352.034 | 0.22         | 3.21           | 4.10        | 1.72         | – –                     | – –         | 11.1      | –       | 0.26      | 2.00     | 501            | 0.15        |
| HD 150193   | 55352.047 | 0.59         | 8.83           | 7.66        | 4.93         | 0.87                     | 1.61        | 11.9      | 0.253    | 0.035     | 1.69     | 560            | 0.25        |
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| HD 190073   | 55329.397 | 0.22         | 3.21           | 4.10        | 1.72         | – –                     | – –         | 11.1      | –       | 0.26      | 2.00     | 501            | 0.15        |
| PDS 2       | 55375.418 | 0.16         | 0.627          | 0.358       | – –          | 3.16                     | – –         | 0.015     | 0.015    | 0.001     | 0.015    | 0.02           | 0.03        |

Table 6  Mean errors of measurements presented in Table 5

| Object      | ±\( \sigma(\text{EW}) \) [\( \text{Å} \)] Ca II 8662 | Ca II 8542 | H\( \alpha \) | Na I D | He I 5876 | H\( \beta \) | W(\( H\alpha10\%) \) | \( \Delta D_B \) |
|-------------|-----------------|-------------|-----------|--------|-----------|---------|----------------|-------------|
| HD 97048    | 0.25             | –           | –         | 0.45   | 0.015     | 0.10    | 20            | 0.03        |
| HD 100546   | 0.30             | –           | –         | 0.45   | 0.015     | 0.20    | 20            | 0.03        |
| HD 101412   | 0.20             | –           | –         | 0.30   | 0.05      | 0.10    | 20            | 0.03        |
| HD 135344B  | 0.15             | –           | –         | 0.25   | 0.045     | 0.10    | 20            | 0.03        |
| HD 150193   | 0.25             | 0.07        | 0.10      | 0.25   | 0.04      | 0.10    | 20            | 0.03        |
| HD 176386   | 0.08             | –           | –         | 0.05   | –         | –       | –             | –           |
| HD 190073   | 0.25             | 0.15        | 0.20      | 0.40   | 0.10      | 0.15    | –             | 0.03        |
| PDS 2       | 0.10             | 0.06        | –         | 0.20   | 0.02      | 0.06    | 20            | –           |

5  Accretion rates

5.1 Testing existing empirical calibrations

We computed the \( \dot{M}_{\text{acc}} \) values for HAes in our sample using the measured quantities from Table 4 for indicators specified at the beginning of Sect. 4. HD 176386 was excluded from our analysis for the reason discussed in Sect. 4.1. The formal calculation of \( \dot{M}_{\text{acc}} \) of this object leads to \( \log \dot{M}_{\text{acc}} = -8.70 \) for Br\( \gamma \), \ -9.03
for Brβ, and −9.63 for Hα. Other indicators are absent in all observed spectra of the star or too weak to be measured precisely.

The indicator \( \Delta D_B \) was used only for six objects with \( \dot{M}_{\text{acc}} > 10^{-8} \text{M}_\odot/\text{yr} \). As was concluded in Donehew & Brittain (2011), this diagnostics is inefficient for \( \Delta D_B < 0.1'' \). The width \( \Delta V(\text{H}\alpha10\%) \) cannot be measured correctly for HD 190073 because a strong P Cyg-type structure overlaps the blue wing of the emission Hα profile. Taking into account that such features are not common to all HAes, there is a doubt if it is reasonable to use them at all. Still, it is of interest to compare the \( \dot{M}_{\text{acc}} \) estimates derived from these diagnostics with values obtained from other indicators.

5.1.1 Criterion of applicability

Our aim was to choose diagnostics that lead to results consistent with each other, where the basic indicator \( \Delta D_B \) is the only tracer connected with \( \dot{M}_{\text{acc}} \) directly on the basis of model calculations. We introduce the criterion of applicability of a diagnostic “\( X \)” for measuring the \( \dot{M}_{\text{acc}} \) of HAes as follows: the indicator is applicable if the mean value of \( < \log \dot{M}_{\text{acc}}(\Delta D_B) - \log \dot{M}_{\text{acc}}(X) > \) (further in the text as “\( \Delta D_B - X \)” ) is consistent with zero within the 1 \( \sigma_m \) errors, calculated for our sample. Here \( \sigma_m = \sqrt{\frac{1}{n} - \bar{X}} \) is the standard error of the mean value and \( n \) is the number of objects.

As a result, we identified a number of diagnostics satisfying this criterion (see Table 7). A small disagreement is present in the Paβ (Dahm 2008) indicator: \( \pm 0.30 \pm 0.09 \) dex, and in the Ca II \( \lambda 8542(L) \) (Dahm 2008) indicator. In any case, the mean spread of all these differences type “\( \Delta D_B - X \)” is of the order of \( \pm 0.10 \pm 0.15 \) dex. It is remarkable that if the lines Na I D and IR Ca II \( \lambda 8542 \) are clearly visible in the spectrum of an object, the \( \dot{M}_{\text{acc}} \) estimates derived from these indicators are in a satisfactory agreement with values obtained from other spectral diagnostics.

The three remaining indicators of \( \dot{M}_{\text{acc}} \) that make use of F-type calibrations (Ca II \( \lambda 8542 \), Ca II \( \lambda 8662 \), and He I \( \lambda 8576 \)) were additionally examined because the “\( \Delta D_B - X \)” values of two Ca II (F) indicators demonstrate differences at larger amplitudes than all other tracers presented in Table 7 \( (+0.44 \pm 0.32 \) dex and \( -0.28 \pm 0.35 \) dex for Ca II \( \lambda 8542 \) (F) and Ca II \( \lambda 8662 \) (F), respectively). Further, the He I \( \lambda 8576 \) tracer shows a notable systematic shift of “\( \Delta D_B - X \)” = \( +0.69 \pm 0.19 \) dex.

5.1.2 Testing F-calibrations

The analysis of \( \dot{M}_{\text{acc}} \) estimates that were obtained using these diagnostics demonstrate considerable systematic inconsistencies between each other as well as between F- and L-type calibrations for the same lines. These discrepancies are illustrated in Fig. 8.

The estimates derived from Ca II \( \lambda 8542 \) are systematically \( 0.75 \pm 0.03 \) dex higher than those from Ca II \( \lambda 8662 \) (top left panel). These two lines originate in the same circumstellar region, and we conclude that conditions in this region in HAes are different from those in TTS for which these two F-calibrations have been introduced.

The mean difference between values obtained from \( \lambda 8542(F) \) and \( \lambda 8542(L) \) is \( +0.90 \) dex for HD 190073 and \( +0.40 \) dex for HD 150193 (top right panel). A similar picture is observed if we compare the \( \dot{M}_{\text{acc}} \) derived with He I \( \lambda 8576 \) (F) and He I \( \lambda 8576 \) (L) calibrations. The values obtained with the F-type calibrations are systematically higher with the spread lying between 0.15 and 0.70 dex for different objects. Each star has its own offset which is the same at all observed epochs pointing at some systematic factor. Based upon the results of this test, we assume that the L-type calibrations are best suitable for the HAes in our sample. We suggest that the observed systematic differences between the values derived with the F- and L-type calibrations are likely to be caused by an insufficient accuracy of the \( c \) and \( d \) coefficients in the F-type calibrations applied to HAes, as well as by uncertainties in \( M_* \) and \( R_* \) of the targets, used for the calculation of \( L_{\text{acc}} \).

5.1.3 Other factors leading to systematic effects

The uncertainties in stellar parameters (especially in the stellar radius \( R_* \)) can be a source of considerable systematic errors in the \( \dot{M}_{\text{acc}} \) determination. As an example, we computed the mass accretion rate of HD 190073 separately with two different values of its stellar radius: \( R_*/R_\odot = 3.6 \) (Catala et al. 2007) and the more recent estimate \( R_*/R_\odot = 2.1 \) (Montesinos et al. 2009). In the first case, we have obtained a mean value of the accretion rate \( -6.23 \), with a spread at the level of standard deviation of \( \pm 0.17 \) dex using eight spectral indicators derived from luminosities of the emission lines (L-type). This value strongly deviates from...
Table 7  List of diagnostics satisfying our criterion of applicability (besides of $\Delta D_B$). The error of the mean value $\sigma_m$ indicates the accuracy of determination.

| Diagnostic “X” | Reference | Mean value of “$\Delta D_B - X$” [dex] | Number of objects |
|----------------|-----------|----------------------------------------|-------------------|
| Brγ(L)         | Donehew & Brittain (2011) | +0.19 ± 0.15                          | 6                 |
| Paβ(L)         | Dahm (2008)     | +0.30 ± 0.09                          | 6                 |
| Paγ(L)         | Dahm (2008)     | +0.11 ± 0.15                          | 6                 |
| Hα(L)          | Dahm (2008)     | −0.02 ± 0.12                          | 6                 |
| Hβ(L)          | Dahm (2008)     | −0.13 ± 0.09                          | 6                 |
| He I λ5876(L)  | Fang et al. (2009) | +0.17 ± 0.17                          | 6                 |
| Na I D(L)      | Herczeg & Hillenbrand (2008) | +0.10 ± 0.16                          | 3                 |
| Ca II λ8542(L) | Dahm (2008)     | −0.22 ± 0.07                          | 2                 |
| He I10%        | White & Basri (2003) | +0.08 ± 0.10                          | 5                 |

Fig. 8  Illustration of systematic discrepancies between the results of the log $\dot{M}_{\text{acc}}$ [dex] determination obtained from comparing three pairs of diagnostics: Ca II λ8542 (F) vs Ca II λ8662 (F) (left top panel); Ca II λ8542 (L) vs Ca II λ8542 (F) (right top panel); and He I λ5876) (L) vs He I λ5876) (F) (bottom panel). The dashed lines indicate the position of equality of abscissa and ordinate. Results for different objects are marked by different symbols. In the upper panels the filled circles mark the six measurements for HD 190073, while open circles refer to measurements obtained for HD 150193. Ca line diagnostics can be used only for these two stars. In the bottom panel HD 97048 is marked by the open circle, HD 100546 by the open triangle pointing downwards, HD 101412 by the open triangle pointing upwards, HD 135344B by filled triangles pointing upwards, HD 150193 by open squares, HD 190073 by filled circles, and PDS2 by filled triangles pointing downwards.
the −7.14 obtained from $\Delta D_B$ which, according to Fig. 10, is independent of $M_*$ and $R_*$ of a star. Using $R_* = 2.1 R_\odot$ leads to the mean value $−7.15 ± 0.17$, that is in good agreement with the estimation from $\Delta D_B$.

Applying different empiric relations between Paβ and Paγ luminosity, respectively, and the mass accretion rates that have been derived for the TTS by various authors to our sample, we find considerable and systematic discrepancies. This is demonstrated in Fig. 9 where we show the $\dot{M}_{\text{acc}}$ values for the targets that result from the calibrations by Gatti et al. (2008) versus those from Dahm (2008). As can be seen in Fig. 9, the difference between the two calibrations is practically absent for small $\log \dot{M}_{\text{acc}}$ (less than −8.0) and becomes up to 0.40 dex for $\log \dot{M}_{\text{acc}} > −7$. We have to mention that Gatti et al. (2008) have calibrated their relation over the $\log \dot{M}_{\text{acc}}$ range from −9.6 to −8.2 ($M_*$ range 0.1−1.0 $M_\odot$) and Dahm (2008) from −8.7 to −7.2 ($M_*$ range 0.5−2.0 $M_\odot$), that is more similar to values $\dot{M}_{\text{acc}}$ and $M_*$ for our targets. Therefore, we used in our study the diagnostics Paβ(L) and Paγ(L) introduced by Dahm (2008).

5.2 The deduced accretion rates for the targets in our sample

Fig. 10 illustrates the mass accretion rates obtained for seven program stars (except HD 176386) that have been derived using the indicators that passed successfully our applicability test (see Sect. 5.1.1). Because the F-type diagnostics lead to systematic differences in the $\dot{M}_{\text{acc}}$ estimates in comparison with the values obtained with the other indicators (see Sect. 5.1.2 and Fig. 8), we did not include the values of $\dot{M}_{\text{acc}}$ derived from F-type indicators in the calculations of the mean values and the uncertainties of the mass accretion rates. As a result, only L-type indicators were used (except Hα 10% and $\Delta D_B$) and, further in the text, we dropped the abbreviation (L) in the names of the indicators.

The mean values of $\dot{M}_{\text{acc}}$ over the observed epochs and over several diagnostics as well as their standard deviations are presented in Table 8 where estimates from Garcia Lopez et al. (2006) and Mendigutia et al. (2011a) are also given for comparison. One can see that, in spite of the fact that the observations of Garcia Lopez et al. (2006) were carried out in 2004, the discrepancies are within the limits of the standard deviations obtained in both works. On the other hand, the results obtained in Mendigutia et al. (2011a) demonstrate significant differences with our estimates. The difference is +0.94 dex for HD 150193 and +2.15 dex for HD 190073. Using data on $A_V$, $L(H\alpha)$, and $L(Br\gamma)$ presented in Mendigutia et al. (2011a) and from our paper, we estimated contributions of various factors forming the total discrepancy. These contributions are: +0.2 − 0.3 dex follow from differences in fluxes measured on different dates, +0.4 − 0.5 dex result from distinctions in $\dot{M}_{\text{acc}}$ calibrations (for both HD 150193 and HD 190073). The contributions from differences in adopted stellar parameters $M_*$ and $R_*$ are +0.2 dex and +1.5 dex for HD 150193 and HD 190073, respectively. Such a significant discrepancy for HD 190073 is mainly related to the very large value of the stellar radius, $R_* = 8 R_\odot$, used in Mendigutia et al. (2011a). This radius is a factor of four lower than the value adopted in our work (Table 2).

We note that the existence of a HAEBe with $T_{\text{eff}}$ less than 10 000 K, $M_*= 5.1 M_\odot$, and $R_* = 8 R_\odot$ (Mendigutia et al. 2011a) is not consistent with the PMS evolution model by Palla & Stahler (1993). According to their model, a PMS object with the initial $M_* = 5 M_\odot$ starts to become visible at the birthline with $T_{\text{eff}}$ of about 12 000 K, and later, closer to the end of the PMS stage near the main sequence, its effective temperature will become about 16 000 K. With the age of about 10$^6$ years given for HD 190073 in Hubrig et al. (2009) and Mendigutia et al. (2011a), the star has to be located in the H-R diagram already near the main sequence, and with $M_* = 2 − 3 M_\odot$ (Table 5), its stellar radius $R_*$ cannot be as large as $8 R_\odot$. As was discussed in Sect. 2.2, the $\dot{M}_{\text{acc}}$ calibrations derived in Mendigutia et al. (2011a) lead to very large values of the mass accretion rates for a considerable part of their targets, which are not in agreement with the masses of accretion disks ($0.01 − 0.1 M_\odot$) and the length in time of the PMS stage for HAes ($10^5 − 10^6$ years). They demonstrate also significant differences with the results of Donehew & Brittain (2011) and Garcia Lopez et al. (2006).

Our $\dot{M}_{\text{acc}}$ determination is based on:

- consistency of values derived from several independent calibrations,
- agreement of our estimates with those obtained in Garcia Lopez et al. (2006), and
- accordance of our results with the predictions of the PMS evolution model by Palla & Stahler (1993).

Clearly, for HAEs with accretion disk masses of the order of 0.01 − 0.1 $M_\odot$ and a length of the PMS stage of HAes of $10^5 − 10^6$ years, the expected mean accretion rates during the PMS stage are of the order of $10^{-6} − 10^{-8} M_\odot$/yr. These expected rates are similar to the values obtained in our work.

Comparing the accuracy of our estimates with the results obtained by other authors, we can see that the differences are not significant (see Table 9).

6 Variability of the mass accretion rates

Four out of seven of our targets were observed on more than one occasion (from two to eight). As was discussed in the previous sections, all emission lines chosen as indicators of $\dot{M}_{\text{acc}}$ are formed not only in the magnetosphere, but also in the disk and in the wind, and their contribution to the entire emission is different for each line. Investigating a temporal behaviour of $\dot{M}_{\text{acc}}$ using various spectral diagnostics, we have to be certain that a variability revealed from these indicators relates to a change of $\dot{M}_{\text{acc}}$ and not to variations of physical and kinematical conditions in the disk and
Fig. 10 Mass accretion rates of the targets determined from different indicators, as labeled. The solid line shows the mean $\log M_{\text{acc}}$ value and the dashed lines indicate the region of $\pm 1\sigma$ uncertainty where $\sigma$ is the standard deviation.

Table 8 Mean mass accretion rates. The error is given as the standard deviation $\sigma$. Literature data on accretion rates are taken from Garcia Lopez et al. (2006) and Mendigutia et al. (2011a).

| Object     | Number of indicators | Number of observations | $\log M_{\text{acc}}$ [M$_\odot$/yr] | $\log M_{\text{acc}}$ [M$_\odot$/yr] | $\log M_{\text{acc}}$ [M$_\odot$/yr] |
|------------|----------------------|------------------------|--------------------------------------|--------------------------------------|--------------------------------------|
| HD 97048   | 8                    | 1                      | $-6.80 \pm 0.38$                     |                                     | $-7.17$                              |
| HD 100546  | 8                    | 1                      | $-7.23 \pm 0.13$                     |                                     |                                     |
| HD 101412  | 9                    | 1                      | $-7.04 \pm 0.15$                     |                                     |                                     |
| HD 135344B | 8                    | 3                      | $-7.69 \pm 0.30$                     |                                     | $-8.27$                              |
| HD 150193  | 10                   | 2                      | $-7.06 \pm 0.20$                     | $-7.29$                             |                                     |
| HD 190073  | 9                    | 6                      | $-7.15 \pm 0.16$                     |                                     | $-6.12$                              |
| PDS 2      | 7                    | 8                      | $-8.68 \pm 0.36$                     |                                     | $-5.00$                              |
the wind. For example, the emission in such lines as Hα, Hβ, and DNa I originates predominantly in the wind, while that in the He I line and near the Balmer jump in the high-temperature region close to the stellar surface (Pogodin et al. 2012). If the variations take place particularly in the outer circumstellar envelope, they are reflected differently in each of these lines. As a result, the amplitudes and even the character of the measured $\dot{M}_{\text{acc}}$ variability derived from different spectral calibrations can be different too. In such a situation, the revealed changes of $\dot{M}_{\text{acc}}$ must be considered as an artifact.

6.1 Temporal behaviour of individual accretion diagnostics

Fig. 11 illustrates variations of $\dot{M}_{\text{acc}}$ of the two targets HD 190073 and PDS 2 with the largest number of observations determined from all indicators. Since the amplitudes of the variability are smaller than the discrepancies between the values obtained from different diagnostics, we analyzed not the values themselves but the residuals relative to the mean value for each indicator. As one can see in Fig. 11, both objects demonstrate variability which: (a) correlates in a majority of the indicators, and (b) shows amplitudes of the variations that are very similar for different diagnostics. It could be intrinsic $\dot{M}_{\text{acc}}$ variability but also some geometric effect that affects in the same way our view of all the different emitting regions.

However, one indicator exists for both objects showing a distinct temporal behaviour of $\dot{M}_{\text{acc}}$: it is He I $\lambda$5876 for HD 190073 and Hα10% for PDS 2. The character of the variability of the $\dot{M}_{\text{acc}}$ value from the He I line in the spectrum of HD 190073 is similar in comparison with other indicators, but shows a much larger amplitude. In the case of PDS 2, the estimations derived from the Hα10% width strongly differ from all others. A possible cause for this can be related to the existence of small-amplitude variations taking place in the circumstellar gas outside the magnetosphere (the outer disk and the wind) which may influence the full width of the profile but might be insufficient to have a significant effect on the EW of Hα. This could explain a rather small distorting influence of the non-magnetospheric variability onto the $\dot{M}_{\text{acc}}$ derived from the Hα(L) diagnostic and a rather significant – onto that obtained with the Hα10% indicator.

The other two stars with more than one available spectrum, HD 150193 and HD 135344B, demonstrate the temporal behaviour of $\dot{M}_{\text{acc}}$ derived from all calibrations that is similar to the case of HD 190073. The only exception is the He I $\lambda$5876 diagnostic. Variability of $\dot{M}_{\text{acc}}$ values derived from this indicator is quite different from that obtained from all other indicators. Presently, no strict model for the formation of the He I $\lambda$5876 line in the high-temperature circumstellar gas around HAeBes exists. We can only assume that a multi-component variability is likely taking place that is connected with a change of geometric configuration, optical thickness and emissivity of the gaseous flows accreted onto the star inside the magnetosphere. But this has a rather small amplitude and the mean value derived from He I differs not so strongly from all the others estimates (see Fig. 10).

6.2 Global accretion variations

As a measure for the overall accretion variability, Fig. 12 illustrates the temporal behaviour of the mean values of $\dot{M}_{\text{acc}}$ for the four targets derived from several diagnostics. The error of the mean value $\sigma_m=\sigma/\sqrt{n-1}$ (where $\sigma$ is the standard deviation and $n$ is the number of the used diagnostics) is shown by vertical bars. Indicators showing a temporal behaviour of $\dot{M}_{\text{acc}}$, different from all others were excluded from the calculation of the mean values and their errors (these are He I $\lambda$5876 for HD 190073, HD 135344B, and HD 150193, and Hα10% for PDS 2).

HD 190073 demonstrates variability on the timescale of tens of days with a spread of the values of about 0.1 dex with an accuracy at $1\sigma_m$ level of ±0.01 dex. Short-term variations (from night to night) are undoubtedly found in PDS 2. The observations of HD 135344B and HD 150193 were carried out only on two to three different nights, separated by several to tens of days. Thus, it is impossible to examine the cause of the variability.

An effective way to analyze the amplitudes and the timescale of accretion variability was suggested by Nguyen et al. (2009), where a change in amplitude versus the time interval between observations was constructed for 40 classical TTS. The authors sampled the timescale range from hours to months and concluded the maximum extent of the $\dot{M}_{\text{acc}}$ variability is reached after a few days and amplitudes of variations are as a rule not more than 0.5 dex. This result has been confirmed by other authors. Costigan et al. (2012) bolster it with a sample of 10 accreting TTS observed in

### Table 9

Mass accretion rate determinations using several diagnostics presented in this work and the results published in the literature.

| Paper                  | Range of log $\dot{M}_{\text{acc}}$ [$M_\odot$/yr] | Range of individual deviations [dex] | Mean standard deviation $\sigma$ [dex] | Error of the mean value $\sigma_m$ [dex] |
|------------------------|-----------------------------------------------------|--------------------------------------|---------------------------------------|----------------------------------------|
| This paper             | $-8.68 \div -6.69$                                  | 0.13 – 0.36                          | 0.25                                  | 0.11                                   |
| Dahm (2008)            | $-8.72 \div -7.20$                                  | 0.1 – 1.4                            | 0.30                                  | 0.12                                   |
| Rigliaco et al. (2011) | $-9.86$                                             | 0.4 – 1.4                            | 0.45                                  | 0.14                                   |
Fig. 11 Temporal behaviour of \( \log \dot{M}_{\text{acc}} \) in HD 190073 and PDS 2 obtained from different indicators, constructed relative to the mean value for each indicator. Thin solid lines join neighbour points for each indicator for better illustration. The temporal behaviour of the residuals \( \log \dot{M}_{\text{acc}} \) obtained from the indicator He I (for HD 190073, left panel) and from the indicator H\( \alpha \)10% (for PDS 2, right panel) are marked by open symbols and thin dashed lines. Dotted lines indicate positions of zero for each indicator.

Due to the insufficient number of observations, it is not possible to analyze the \( M_{\text{acc}} \) variability on different timescales for the other three targets. The only conclusion can be made that a variability has been revealed on timescales from several days to tens of days. In the case of HD 190073 we can derive the rotation period \( P = 9.1 \sin i \) days using data from Table 2. In Hubrig et al. (2009) the upper limit of \( i \) is given as \( i \leq 40^\circ \). A P Cyg-type of the H\( \alpha \) profile (Fig. 3) evidences that this object has to be of an intermediate orientation relative to an observer (Grinin & Rostopchina 1996). With the expected value of \( i \sim 40^\circ \) we estimate \( P \sim 6 \) days. Since no observations separated by short time intervals have been carried out for this target so far, we cannot examine a presence of rotational modulations of the line profiles in the spectrum of HD 190073.

Another plausible alternative could be that the variability is the result of stochastic variations in the accretion rate through the disk, as one might expect if the magneto-rotational instability is the primary source of viscosity in accretion disks. More spectral data obtained on different timescales (from days to tens of days) are needed to rigorously test the character of the variability of the Herbig Ae/Be stars.

7 Conclusions

We present results of observations of a sample of eight magnetic Herbig Ae stars obtained with the X-shooter spectro-
Fig. 12 Temporal behaviour of the mean values of the residual log $\dot{M}_{\text{acc}}$ of four targets. The errors at $\pm 1 \sigma_m$, where $\sigma_m$ is the error of the mean, are marked by bars. The values are obtained using all indicators for each object (see Fig. 8) except He I (for HD 190073, HD 135344B, and HD 150193) and H$\alpha$10% (for PDS 2).

An important cause of errors is the uncertainty in determinations of $M_*$ and especially $R_*$ for a number of targets when the indicators are based on luminosity determination. For objects with small accretion rate values ($\log \dot{M}_{\text{acc}} < -8.5$) errors of measurements of EWs and widths of emission profiles become an additional factor for the measurement accuracy.

Our investigation of variability of $\dot{M}_{\text{acc}}$ has shown that all four Herbig Ae stars observed on more than one occasion demonstrate a change of $\dot{M}_{\text{acc}}$ of about $0.10 \div 0.40$ dex on the timescale of tens of days. One object observed on a short timescale, PDS 2, shows also night-to-night variations with an amplitude of up to $0.30$ dex. This variability might be related to the modulation of spectral parameters by a rotating magnetosphere. In the case of even longer-term variability ($\tau \approx$ tens of days) it remains unclear whether this variability is connected with a change in the accretion regime or is a result of rotational modulation. We note that due to the small number of observations of targets with determined magnetic/rotation periods, not much can currently be concluded on the role of magnetic fields in the dynamics of the accretion processes. Future multi-epoch observations of magnetic HAes with X-shooter will be extremely useful to better understand the nature and variability of $\dot{M}_{\text{acc}}$ spectral diagnostics to constrain more realistically magnetospheric accretion models for Herbig Ae stars.
Fig. 13 Amplitude of variations of the mean $\log M_{acc}$ of four targets as a function of time elapsed between the observations. Upper panel: Results for HD 190073 (filled symbols) and HD 135344B (open symbols). Bottom panel: Results for PDS 2 (filled symbols) and HD 150193 (open symbols).

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