1. INTRODUCTION

Cometary ices have undergone little or no processing since their formation in the solar nebula, and thus they represent an important clue in understanding the conditions in the early solar system. However, this interpretation is complicated by radial mixings in the protoplanetary disk and the dynamical evolution of comets, expelled to the outer solar system from their original formation site. No pattern for the cometary composition and activity has emerged from the objects studied to date (Biver et al. 2002; Crovisier et al. 2007), and a larger sample is needed. The composition of cometary ices has been studied through spectroscopic observations or in situ measurements. In the analysis of cometary spectra it is important to establish whether the observed chemical species originate in the nucleus (native source) or in the coma (extended source). For this purpose, spatially resolved observations are needed to derive the radial composition of the gas outflow.

Carbon monoxide (CO) is one of the most abundant molecules, and its mixing ratio relative to water varies greatly among comets (Biver et al. 2006; Bockelée-Morvan et al. 2004). As a highly volatile compound with a sublimation temperature of 24 K, CO ice is thought to be a sensitive tracer of the temperature in the environment in which the comets formed. However, its origin as a native compound or a daughter product, and the correlation between the two sources is still not understood. Infrared observations of the CO rovibrational lines conclude that the CO source was mainly native in comet C/1996 B2 (Hyakutake; DiSanti et al. 2003) and significantly extended for comet C/1995 O1 (Hale-Bopp; Disanti et al. 2001). This paper complements previous findings with the first CO radial profiles constructed from ultraviolet spectra. Care must be taken in interpreting the brightness profiles, since saturation close to the nucleus could mimic the presence of an extended source. A comprehensive fluorescence model is developed to derive reliable CO column densities and estimate CO production rates from cometary spectra obtained by the Space Telescope Imaging Spectrograph (STIS) and the Goddard High Resolution Spectrograph (GRHS) on the Hubble Space Telescope (HST). Using the long-slit capabilities of the STIS instrument, this paper offers evidence that the native source of CO is dominant in three long-period comets.

CO ultraviolet fluorescence in cometary spectra was first detected during sounding rocket observations of comet West (C/1975 V1; Feldman & Brune 1976) and has been subsequently observed by IUE (Tozzi et al. 1998), HST (Weaver 1998), FUSE (Feldman et al. 2002), the Hopkins Ultraviolet Telescope on the Astro-1 space shuttle mission (Feldman et al. 1991), and rockets (Woods et al. 1987; Sahnow et al. 1993; McPhate et al. 1999). The most important spectral features of CO in the ultraviolet (UV) are its electronic transitions belonging to the $A-X$, $B-X$, and $C-X$ systems, or to the forbidden Cameron bands. Fluorescence is the main emission mechanism for the $A^1\Pi - X^1\Sigma^+$ system (1300–1900 Å), $C^1\Sigma^+ - X^1\Sigma^+$ system (0–0 band at 1087.9 Å), and $B^1\Sigma^+ - X^1\Sigma^+$ system (0–0 band at 1150.5 Å). The Cameron bands, $a^3\Pi - X^1\Sigma^+$ (1900–2800 Å), are mainly excited by electron impact and photodissociation of CO$_2$ (Weaver et al. 1994). Although observations of the $A-X$, or Fourth Positive Group, of CO in the UV spectra of comets have a long history, their interpretation has been difficult compared to that of the $B-X$ and $C-X$ systems at shorter wavelengths that have been observed more recently.
The spatial information offered by the high-resolution long-slit HST STIS spectra of comets 153P/Ikeya-Zhang (C/2002 C1), C/2001 Q4 (NEAT), and C/2000 WM$_1$ (LINEAR) shows that close to the nucleus the CO emission in the $A$ -- $X$ system is self-absorbed. This follows from the observed change in the relative intensities in various vibrational progressions as the offset from the comet center increases (see § 3). Self-absorption makes it difficult to derive a reliable value for the CO column density in the absence of a model that takes into account optical depth effects.

A detailed model of the $A$ -- $X$ system is needed to track the effects of saturation and self-absorption. We constructed a database containing ~$10^3$ transitions between the first 50 rotational levels of each of the 37 vibrational levels of $X^1\Sigma^+$ and the 23 vibrational levels of $A^1\Pi$, taking into account the energy shifts and mixings of the transition probabilities due to interactions between the different parity sublevels (Le Floch et al. 1987; Morton & Noreau 1994). Using a simple approximation for fluorescence in subordinate lines (Liu & Dalgarno 1996), expanded with a comprehensive treatment of self-absorption, our model offers an excellent fit to the data. The $A$ -- $X$ fluorescence model is used to derive spatial profiles of the CO column density for the three comets observed by STIS. We find that the column density profiles are consistent with a dominant native source in all comets. The resulting production rates range between $3.56 \times 10^{25}$ and $1.76 \times 10^{26}$ molecules s$^{-1}$ and are corroborated by the results from high-resolution FUSE observations of C/2001 Q4 (NEAT) and C/2000 WM$_1$ (LINEAR; Weaver et al. 2002). The HST GHRS observations of comet C/1996 B2 (Hyakutake), a comet with a high CO production rate (Biver et al. 1999; DiSanti et al. 2003) and thus strongly affected by large optical depth effects, require the addition of geometrical corrections to the simple plane-parallel atmosphere model in order to reproduce the relative line strengths.

The HST observations are described in § 2. Details about the model and its simplifying assumptions are found in § 3. Detailed data analysis follows in § 4, and a discussion of the results is given in § 5. We conclude with a summary in § 6.

## 2. Observations

The HST observations are summarized in Table 1. Comets Ikeya-Zhang and C/2000 WM$_1$ (LINEAR) were observed near times of high solar activity, while C/2001 Q4 (NEAT) and Hyakutake were observed closer to solar minimum. The HST STIS observations used the G140L grating and the 52X0.2 aperture (25$''$ x 0.2$''$ for L-mode MAMA), resulting in a spatial resolution of ~0.1$''$ and a spectral resolution of ~4 Å in the wavelength range 1150--1730 Å. The STIS instrument performance is described in Kimble et al. (1998) and Woodgate et al. (1998). The HST GHRS (Heap et al. 1995) observations of comet C/1996 B2 (Hyakutake), covering the 1290--1590 Å bandpass with a spectral resolution of ~4 Å, do not provide spatial information. The GHRS Large Science Aperture (LSA) translates into a 1.74$''$ x 1.74$''$ post-COSTAR projected area. Given that FUSE and HST STIS observations of comets C/2000 WM$_1$ (LINEAR) and C/2001 Q4 (NEAT) were nearly simultaneous, we choose to include in Table 1 a summary of FUSE observations for completeness (Feldman et al. 2002; Weaver et al. 2002).

## 3. Fluorescence Model for the $A^1\Pi$ -- $X^1\Sigma^+$ System of CO

The Fourth Positive Group of CO, $A^1\Pi$ -- $X^1\Sigma^+$, has non-negligible overlap integrals for most of the nondiagonal vibrational transitions. While for the $C$ -- $X$ and $B$ -- $X$ systems it is enough to model just the (0--0) and (1--0) bands, for the $A$ -- $X$ system we must take into account all bands connecting all 37 vibrational levels of the $X^1\Sigma^+$ state with all 23 vibrational levels of $A^1\Pi$. This requires solving the system of diatomic equations for all levels of the CO molecule.

### Table 1

| Data Set | Date and Time (UT) | Exposure Time (s) | $r_h$ (AU) | $r_v$ (km s$^{-1}$) | $\Delta$ (AU) | Instrument | Mode |
|----------|-------------------|-------------------|------------|-------------------|-------------|------------|--|
| C/1996 B2 (Hyakutake) | | | |
| Z35FN602T | 1996 Apr 1 13:09:00 | 1305.600 | 0.885 | -38.51 | 0.259 | HST GHRS | G140L; 2.0 |
| Z35FN604T | 1996 Apr 1 13:39:00 | 217.600 | 0.885 | -38.51 | 0.260 | HST GHRS | G140L; 2.0 |
| Z35FN702T | 1996 Apr 1 14:44:00 | 1305.600 | 0.884 | -38.53 | 0.262 | HST GHRS | G140L; 2.0 |
| Z35FN704T | 1996 Apr 1 15:14:00 | 217.600 | 0.883 | -38.53 | 0.262 | HST GHRS | G140L; 2.0 |
| C/2000 WM$_1$ (LINEAR) | | | |
| B0500301000 to B0502401000 | 2001 Dec 10 00:45:00 | 1800.198 | 1.083 | -38.53 | 0.262 | HST GHRS | G140L; 2.0 |
| O6GR11010 | 2001 Dec 9 23:09:00 | 1305.600 | 0.884 | -38.53 | 0.262 | HST GHRS | G140L; 2.0 |
| O6GR03010 | 2000 Dec 9 23:09:00 | 1305.600 | 0.884 | -38.53 | 0.262 | HST GHRS | G140L; 2.0 |
| O6GR11010 | 2000 Dec 9 23:09:00 | 1305.600 | 0.884 | -38.53 | 0.262 | HST GHRS | G140L; 2.0 |
| 153P/Ikeya-Zhang | | | |
| O8FY01010 | 2002 Apr 07:28:00 | 1890.199 | 0.887 | 29.07 | 0.426 | HST STIS | G140L; 52 x 2.0 |
| O8FY01010 | 2002 Apr 10:40:00 | 1890.199 | 0.899 | 29.08 | 0.426 | HST STIS | G140L; 52 x 2.0 |
| O8FY01010 | 2002 Apr 10:40:00 | 1890.199 | 0.899 | 29.08 | 0.426 | HST STIS | G140L; 52 x 2.0 |
| C/2001 Q4 (NEAT) | | | |
| E139010010 | 2004 Apr 24 00:39:00 | 68.282 | 1.030 | -10.80 | 0.510 | FUSE | LWRS |
| O8VK04010 | 2004 Apr 25 20:03:00 | 1890.199 | 0.899 | 29.08 | 0.426 | HST STIS | G140L; 52 x 2.0 |
| O8VK04010 | 2004 Apr 26 00:49:00 | 1890.199 | 0.899 | 29.08 | 0.426 | HST STIS | G140L; 52 x 2.0 |
| O8VK04010 | 2004 Apr 26 00:49:00 | 1890.199 | 0.899 | 29.08 | 0.426 | HST STIS | G140L; 52 x 2.0 |
the $A^1\Pi$ state. Using only the first 50 rotational levels, which should be sufficient for typical physical conditions in cometary comae, the final database contains almost $10^5$ transitions. The latest values for the parameters of the $A^1\Pi$ (Kurucz 1976; Morton & Noreau 1994; Le Floch 1991) and $X^1\Sigma^+$ states (George et al. 1994) were used to derive the energy levels and transition wavelengths. Following Morton & Noreau (1994) the rotational transition probabilities $A_{J',J}$ were obtained from the band transition probabilities $A_{\nu',\nu}$ (Beegle et al. 1999; Eidesberg et al. 1999; Borges et al. 2001; Kirby & Cooper 1989). Even though there are three rotational branches ($P$, $Q$, and $R$; $\Delta J = 0$, $\pm 1$) possible from the same rotational level of the $A^1\Pi$ state ($J'$), due to the parity selection rule, the excitation rates and branching ratios for the $P$ and $R$ branches do not mix with those for the $Q$ branch. Each $J'$ level of the $A^1\Pi$ state is split into two opposite parity sublevels ($\Lambda$ doubling) that have different energies and interactions with neighboring levels of other electronic states (Morton & Noreau 1994). The energy level shifts and the changes in the lifetimes due to these interactions were taken into account when available (Morton & Noreau 1994; Kittrell et al. 1993; Le Floch et al. 1987).

In a first approximation, the fluorescent emission in the Fourth Positive Group of CO is modeled following the prescription from Liu & Dalgarno (1996). This model successfully accounts for saturation in the absorption of the exciting radiation, using a Voigt profile with line overlap. A quick look at the data, however, shows that this approach only partially accounts for optical depth effects, and we must extend the model to include self-absorption in the fluorescent cascade. The STIS spectra of comet 153P/Ikeya-Zhang shown in Figure 1 are averaged over effective apertures of increasing size, centered on the comet nucleus. Self-absorption is easily recognized by noting that the bands connecting to the ground vibrational level ($0^0$, $1^0$, $2^0$, $3^0$, $4^0$, $5^0$) show little decrease in brightness as we increase the integrated slit area and the average column density becomes lower, while the other bands from the progressions originating in the same upper level (see the unblended $0^1$, $1^1$, $2^1$, $3^1$, $4^1$, and $5^1$) decrease strongly in intensity, so that at low column densities the relative line ratios are in agreement with the optically thin limit.

The treatment of self-absorption is made under the assumption of local thermodynamic equilibrium (LTE), using the fact that the photon mean free path corresponds to an optical depth $\tau$ of one. Given that the gas temperature in the coma is less than 100 K for our observations, under LTE conditions we need only to take into account absorption out of the lowest (0) vibrational level of $X^1\Sigma^+$. Unlike those of $\text{H}_2$, for example, the rovibrational levels of the ground state of CO have short lifetimes, making LTE a good approximation. The lines connecting to $\nu' = 0$ for which $\tau$ at line center is greater than one are considered self-absorbed. Only an effective column density $N_0/\sigma_{\text{line-center}}$ smaller than the total absorbing column $N$, will contribute to the observed emission in a self-absorbed transition $i-j$. The excess brightness, due to absorption by the total column density $N$, is then redistributed among the optically thin lines originating in the same upper level $i$ according to the branching ratios, as seen in Figure 2.

The model is constructed in the approximation of a uniform plane-parallel atmosphere. This approximation is of limited use, due to the spherical symmetry of the comet atmosphere and the Sun-comet-Earth geometry. The breakdown is more apparent for larger column densities, such as in the case of comet C/1996 B2 (Hyakutake), discussed in § 4.2. In the optically thin regime, valid at larger distances from the comet nucleus, as well as for unresolved objects, such geometrical effects are negligible. Introducing geometrical corrections in our model can be done by allowing the column density entering the absorption step to be different from the column density used for emission and self-absorption. This procedure accounts for the fact that the projected column density along the line of sight is larger than the column density toward the exciting radiation source. This difference leads both to the decrease in line saturation and to the enhancement of
The only free parameter of the model is the column density, which is varied over a grid of values until the best fit is found. All model parameters used for each comet are summarized in Table 2. The values for the rotational temperature and Doppler b parameter were chosen in agreement with infrared and radio measurements, when available. The Doppler b parameter is given by

\[ b = \left( t_{\text{therm}} + t_{\text{flow}} \right)^{1/2} \approx t_{\text{flow}}, \]

where \( t_{\text{flow}} \) is the source of nonthermal line broadening, and the thermal velocities are comparable to the uncertainties in \( t_{\text{flow}} \). Molecular lines of water and other molecules are well resolved by radio observations (usually to better than 0.1 km s\(^{-1}\)), and have a Gaussian profile, reflecting the symmetric outflow of the gas (Lecacheux et al. 2003; Wouterloot et al. 1998; Biver et al. 1999). Our observations have too low a resolution to be used to determine the Doppler line widths directly, so we rely on the published values for the expansion velocity of the gas, derived from the line profiles after correcting for thermal and instrumental broadening.

4. DATA ANALYSIS

Although the brightness of the \( A - X \) bands for the same comet should differ slightly from one observation to another due to varying comet heliocentric and geocentric distances, and due to possible periodic variations in the volatile vaporization rate, this variation is within the error bars of the observations and the use of an averaged STIS spectrum for each comet in order to improve the signal-to-noise ratio is warranted. We also select only the data sets that do not show significant deviations in background and intensity from one another.

High-resolution solar spectra from the Ultraviolet Spectrometer Polorimeter Experiment on the Solar Maximum Mission (Tandberg-Hanssen et al. 1981) were used, scaled to match the solar activity at the time of comet observations as deduced from Upper Atmosphere Research Satellite (UARS) Solar Stellar Irradiance Comparison Experiment (SOLSTICE) solar flux measurements (Rottman et al. 2001). The solar spectrum is shifted according to the comet motion relative to the Sun (Swings effect; Dymond et al. 1989). For the solar \( \mathrm{H} \alpha \) lines \( \lambda 1561.0 \) and \( \lambda 1657.6 \) multiplets. Using the spatial information available, we derive the CO spatial profile for each comet by fitting spectra extracted from regions of 1.5 arcmin width at increasing offsets from the comet center. The selected regions sample areas of varying column density, for which our model estimates an average value. The comet nucleus is identified with the center of brightness, located at zero in Figure 3. The innermost region extracted, centered on the nucleus, is noisier due to the small integrated area. For intermediate regions the signal is better, due to the averaging of two regions, symmetric about the comet center. For each region, the background subtraction is performed by fitting a quadratic polynomial to selected points from feature-free intervals. These points were selected such that the outliers are discarded and the resulting polynomial fit is optimal.

The best-fit column density for each region and its standard deviation are derived by minimizing the \( \chi^2 \) statistics, taking into account the errors in background subtraction. The range of column densities obtained for each comet observed by STIS is listed in Table 2, together with the average value for comet C/1996 B2 (Hyakutake). The results are further compared with the Haser native source model (Haser 1957; Opal & Carruthers 1977), with an outflow velocity of \( 0.85 r_b^{0.5} \) km s\(^{-1}\), where \( r_b \) is the comet-Sun distance in AU (Budzien et al. 1994; Biver et al. 1999), and a CO lifetime of \( 2 \times 10^5 \) s. We derive the CO production rate for each comet by a least-squares fit of the Haser model to the radial

### Table 2: Model Parameters

| Comet Name                | \( T_{\text{tot}} \) (K) | Doppler b Parameter (km s\(^{-1}\)) | Solar Activity | \( N_{\text{CO}} \) Range (10\(^{12}\) cm\(^{-2}\)) |
|--------------------------|-------------------------|-----------------------------------|----------------|----------------------------------|
| 153P/Ikeya-Zhang          | 82\(^{a}\)              | 0.91\(^{b}\)                      | Maximum        | 61.3–1.49                        |
| C/2001 Q4 (NEAT)          | 68\(^{b}\)              | 0.79\(^{d}\)                      | Minimum        | 68.4–1.86                        |
| C/2000 WM_{j} (LINEAR)    | 77\(^{e}\)              | 0.72\(^{b}\)                      | Maximum        | 0.935–0.327                      |
| C/1996 B2 (Hyakutake)     | 72\(^{e}\)              | 2.0\(^{f}\)                       | Minimum        | 145                              |

\( ^{a} \) From 74\( r_b^{0.3} \) K dependence, Dello Russo et al. (2004).

\( ^{b} \) Biver et al. (2006).

\( ^{c} \) Temperature of cold component, FUSE observations (see text).

\( ^{d} \) Estimated from outflow velocity 0.85\( r_b^{0.5} \).

\( ^{e} \) From 63\( r_b^{1.0} \) K dependence, DiSanti et al. (2003).

\( ^{f} \) Within the range of values given by Wouterloot et al. (1998).
column density profile. The native source model is integrated over rectangular regions matching the 1.5\arcsec spectral extractions along the STIS slit. The resulting production rates and their magnitude relative to water are listed in Table 3.

4.1.1. 153P/Ikeya-Zhang

The values for the rotational temperature and Doppler b parameter are 82 K and 0.91 km s\(^{-1}\) respectively, derived from infrared and radio measurements (Dello Russo et al. 2004; Biver et al. 2006). Spectra extracted from 1.5\arcsec intervals were fitted using CO column densities ranging from $6.13 \times 10^{15}$ to $1.49 \times 10^{14}$ cm\(^{-2}\), as shown in Figure 4. The data shown are obtained by averaging the first and third STIS observations (Table 1), and the best-fit model is plotted in red. The second observation was not included due to the background mismatch with the other two. The derived values for the CO column density in each region are represented by stars in Figure 5, plotted as a function of the distance from the comet nucleus. The error bars are given by the 1 \(\sigma\) confidence level from the \(\chi^2\) statistics. Fitting to this radial profile a native source model integrated over rectangular regions with the same coverage as the extracted spectra (Opal & Carruthers 1977), we obtain a production rate of \((1.54 \pm 0.09) \times 10^{28}\) molecules s\(^{-1}\). The native source model for the derived production rate is shown as a continuous line in the same figure. The resulting CO production rate relative to water is about $7.2\% \pm 0.4\%$. The water production rate \((2.15 \times 10^{29}\) molecules s\(^{-1}\)) was obtained from a vectorial model fit to an HST STIS observation of the OH (0–0) band made on 2002 April 21 at 12:19 UT.

The residuals after subtracting the fluorescence spectrum for the CO A – X system reflect the contributions of atomic species and allow the first detection of H\(_2\) at wavelengths longward of 1200 Å. The H\(_2\) spectrum consists of the P(1) lines of the Lyman (6-\(v'\)) progression pumped by solar H I Ly\(\beta\). The H\(_2\) lines with \(v' = 1–3\) were first detected in the FUSE spectrum of comet C/2001 A2 (LINEAR; Feldman et al. 2002). The lines at longer wavelengths, including the strongest one in the progression (\(v' = 13\) at 1607.5 Å), remained undetected due to the abundance of CO features and lower resolution of the STIS instrument. The solar maximum value of flux, together with a velocity shift that placed the (6–0) P(1) line at the peak of the line, made it particularly fortuitous to detect the (6–13) P(1) line in comet Ikeya-Zhang. Figure 6 shows the shape and intensity of the solar Ly\(\beta\) at minimum and maximum activity and the Doppler shifts of the H\(_2\) absorption line corresponding to each comet. Although H\(_2\) lines pumped by Ly\(\beta\) were detected in comet C/2001 Q4 (Feldman et al. 2004), due to the large negative Doppler shift and the low solar activity, in the STIS bandpass the signal to noise for the H\(_2\) lines is too low to warrant a detection. A comparison of the residuals for the two comets after subtracting the CO fluorescence model is shown in Figure 7. The STIS spectrum was integrated over a 4\arcsec wide region centered on the comet nucleus, and the CO column density for the two comets \((N = 3.07 \times 10^{12}\) and $3.49 \times 10^{15}$ cm\(^{-2}\), respectively) has been obtained through the \(\chi^2\)

![Image](https://example.com/image.png)

**Figure 3.** HST STIS spectral image of comet 153P/Ikeya-Zhang, showing the line brightness variation along the slit. The spectral region containing the strong geocoronal H I \(\lambda\)1216 and O I \(\lambda\)1302 has been excluded. The carbon multiplets at 1561.0 and 1657.6 Å are relatively constant due to the extended source component.

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### Table 3: Production Rates

| Comet Name                  | \(\dot{Q}_{\text{CO}}\) \(\times 10^{28}\) molecules s\(^{-1}\) | \(\dot{Q}_{\text{CO}}/\dot{Q}_{\text{H}_2}\) (%) | Other \(\dot{Q}_{\text{CO}}\) Measurements \(\times 10^{28}\) molecules s\(^{-1}\) |
|-----------------------------|-------------------------------------------------------------|------------------------------------------------|-----------------------------------------------------------------------------|
| 153P/Ikeya-Zhang            | 1.54 \pm 0.09                                              | 7.2 \pm 0.4\(^a\)                                      | 0.73 \pm 0.16\(^b\)                                                      |
| C/2001 Q4 (NEAT)            | 1.76 \pm 0.16                                              | 8.8 \pm 0.8\(^a\)                                      | 1.36 \pm 0.40\(^c\)                                                      |
| C/2000 WM\(_3\) (LINEAR)    | 0.036 \pm 0.002                                            | 0.44 \pm 0.03\(^f\)                                     | 0.035 \pm 0.003\(^d\)                                                    |
| C/1996 B2 (Hyakutake)       | 4.97 \pm 0.07                                              | 20.9 \pm 0.3\(^h\)                                     | 4.84 \pm 0.58\(^i\)                                                      |

\(^a\) The error bars are given by the 1 \(\sigma\) interval from the \(\chi^2\) statistics. In addition, we estimate that systematics amount to a 15\% uncertainty in the production rates for the STIS observations.

\(^b\) Water production rate of $2.15 \times 10^{29}$ molecules s\(^{-1}\) from HST STIS observations (see text).

\(^c\) Using \(r_{\text{CO}}^{2.1}\) scaling from Biver et al. (2006).

\(^d\) Water production rate of $2.0 \times 10^{29}$ molecules s\(^{-1}\) from HST STIS observations (see text).

\(^e\) Derived in this paper from FUSE observations (§ 4.1.2). The cited value includes only the cold source component.

\(^f\) Water production rate of $8.0 \times 10^{28}$ molecules s\(^{-1}\) used for the FUSE observations (Weaver et al. 2002).

\(^g\) Based on FUSE observations (Weaver et al. 2002), revised for this paper. The cited value includes only the cold source component.

\(^h\) Water production rate of $2.38 \times 10^{29}$ molecules s\(^{-1}\) from Combi et al. (1998).

\(^i\) Value measured on April 1.2 (UT) using JCMT radio telescope (Biver et al. 1999). From the $4.7 \times 10^{29}$ \(r_{\text{CO}}^{2.1}\) molecules s\(^{-1}\) dependence the predicted value is $6.07 \times 10^{28}$ molecules s\(^{-1}\).
minimization procedure. The residuals for comets C/2000 WM1 and Hyakutake (not included in the figure) show no evidence for H$_2$, which can be explained by the large negative Doppler shifts with respect to the solar Ly$\beta$, the low outgassing rate of C/2000 WM1, and the minimum solar activity at the time of Hyakutake observations.

The synthetic H$_2$ fluorescence spectra shown in red in Figure 7 are constructed under the assumption of an H$_2$O photodissociation source for H$_2$. According to the dissociation model, H$_2$ is rotationally hot (100–300 K) and its rotational levels are in statistical equilibrium, with an ortho/para ratio of three, similar to water (Budzien et al. 1994; van Harrevelt & van Hemert 2000;
Bonev et al. (2007). The uncertainty in the exact value for the rotational temperature has little impact on the emerging spectrum, as the population of the \( J = 1 \) level of the ground state does not differ significantly for rotational temperatures ranging from 100 to 300 K. The fluorescence efficiencies, or \( g \)-factors, used in the model were revised from those cited by Feldman et al. (2002) using the solar Ly\( \beta \) profiles obtained over the past solar cycle by Lemaire et al. (2002) and accounting for the comet’s heliocentric velocity. The quality of the data does not warrant a \( \chi^2 \) fit for the \( \text{H}_2 \) fluorescence model, so we restrict ourselves to making rough estimates. Using an \( \text{H}_2 \) column density of \( 1.0 \times 10^{14} \text{ cm}^{-2} \) at a rotational temperature of 200 K we obtain a reasonable agreement with the residuals for 153P/Ikeya-Zhang (Fig. 7, top). As was the case for comet C/2001 A2, this value for the \( \text{H}_2 \) column is consistent, within the rather large uncertainties in both the data and the models, with an \( \text{H}_2 \text{O} \) photodissociation source. Constraining the \( \text{H}_2 \) production helps our understanding of this \( \text{H}_2 \text{O} \) photodissociation channel, for which little laboratory data is available.

Under our choice of parameters, the model for 153P/Ikeya-Zhang predicts about 5.5 R for the \( v'' \) = 7, 9, and 11 lines, at the level of the errors due to noise and \( \text{CO} \) subtraction, leaving only the \( (6-13) \) \( P(1) \) line detectable at a \( \approx 3 \sigma \) level, with 18.4 R. This line is not detected in any other comet from our sample. For comparison, we model the \( \text{H}_2 \) fluorescence for comet C/2001 Q4. We use the same column density of \( \text{H}_2 \), as both C/2001 Q4 and 153P/Ikeya-Zhang have similar water production rates (see notes to Table 3). Due to the low solar activity and the unfavorable Doppler shift, the predicted line intensities are too low to be detected, consistent with the residuals (Fig. 7, bottom). The discrepancies between the \( \text{H}_2 \) models and the residuals are due to the uncertainties in the CO fluorescence model itself, as well as to the noise in the data increasing toward longer wavelengths.

The strongest atomic lines seen in the residuals are also indicated in Figure 7. In addition to those lines usually seen in comets (MePhate et al. 1999), we also note the presence of the \( \text{S} \) \( (1D-1P^o) \) transition at 1666.7 Å. This transition is analogous to \( \text{O} \) \( (1D-1P^o) \) at 1152.2 Å (Feldman et al. 2002) and \( \text{C} \) \( (1D-1P^o) \) at 1930.9 Å (Tozzi et al. 1998). Using a \( g \)-factor of 1.92 \( \times 10^{-5} \) photons atom\(^{-1} \) s\(^{-1} \) at 1 AU, the \( \approx 30 \) R brightness corresponds to an average column density of \( \text{S} \) \( (1D) \) atoms in the aperture of \( 1.24 \times 10^{12} \text{ cm}^{-2} \). Since the lifetime of the metastable \( 1D \) state of sulfur is only 28 s, this requires that within the aperture \( 1D \) atoms be produced at a rate of \( 1.9 \times 10^{28} \text{ s}^{-1} \) (0.09% relative to water). Collisional de-excitation of \( 1D \) near the nucleus would raise this number. The likely source of these atoms is the photodissociation of sulfur-bearing molecules such as \( \text{H}_2 \text{S} \) or \( \text{CS}_2 \), which must be produced at a rate greater than 0.1% relative to water.

4.1.2. C/2001 Q4 (NEAT)

Adjusting the model parameters to the conditions of NEAT Q4 observations, we derive in the manner described for 153P/Ikeya-Zhang a \( \text{CO} \) production rate of \( (1.76 \pm 0.16) \times 10^{29} \text{ molecules s}^{-1} \), or 8.8% \( \pm 0.8% \) relative to water. For the water production rate, we derive a value of \( 2.0 \times 10^{29} \text{ molecules s}^{-1} \) from STIS observations on 2004 April 23 21:39 UT. The CO model was fitted to the average of the first two STIS observations (Table 1). For the third observation the detector background level is much higher than for the other two. The adopted values for the rotational temperature and Doppler \( b \) parameter are listed in Table 2. \text{FUSE} observations of the \( \text{CO} \) \( \text{C} - \text{X} \) Hopfield-Birge system at 1088 Å reveal a band profile consistent with a two-component temperature model (Feldman et al. 2002). The hot component (\( \approx 600 \) K) is described to an extended CO source that is due to the dissociation of \( \text{CO}_2 \) (Feldman et al. 2006b). We use for our model the temperature of the cold component, estimated at \( \approx 68 \) K, characteristic of the native CO source which dominates at the smaller cometocentric distances probed by STIS. Lacking direct radio measurements of the line widths, we use a value of 0.79 km s\(^{-1} \) for the Doppler \( b \) parameter, based on the outflow velocity 0.8R\( \times 0.5 \). The radial profile of the best-fit column densities and the native source model are plotted in Figure 8 as stars and a solid line, respectively.

Using a two-component fit to the \( \text{CO} \) \( \text{C} - \text{X} \) band observed by \text{FUSE} (Feldman et al. 2002), we derive \( \text{CO} \) column densities of \( (1.2 \pm 0.3) \times 10^{14} \text{ cm}^{-2} \) and \( (2.38 \pm 0.08) \times 10^{13} \text{ cm}^{-2} \) for the cold and hot components, respectively, averaged over the entire \( 30'' \times 30'' \) slit. A native source for the cold component requires a production rate of \( (1.36 \pm 0.40) \times 10^{28} \text{ molecules s}^{-1} \), or 6.8% \( \pm 2.0% \) relative to water. The solar flux pumping the \( \text{C} - \text{X} \) fluorescence was based on quiet Sun whole disk data from \text{SOHO SUMER} (Curdt et al. 2001), normalized at wavelengths longward
of 1200 Å to UARS SOLSTICE solar flux measurements appropriate to the solar activity at the time of our observation (Rottman et al. 2001).

4.1.3. C/2000 WM1 (LINEAR)

The CO emission detected by STIS in the observation of comet C/2000 WM1 (LINEAR) was too weak to allow us to repeat the same analysis performed in the case of comets 153P/Ikeya-Zhang and C/2001 Q4 (NEAT). Instead, we chose to integrate the STIS spectrum over increasing widths centered on the nucleus, in order to make use of the stronger signal in the center and to increase the number of contributing pixels. We started with a 4″ wide region which was increased progressively up to 16″. For the rotational temperature we used the 77 K value derived for the cold component using FUSE observations (Weaver et al. 2002), while the Doppler b parameter value of 0.72 km s\(^{-1}\) was chosen to match the radio observations of Biver et al. (2006). All three STIS observations were averaged together to obtain detectable CO emission features. The best-fit column densities over the selected regions are plotted in Figure 9 as a function of the integrated slit width.

![Figure 7](image1.png)

Fig. 7.—Residuals from the spectrum of 153P/Ikeya-Zhang and C/2001 Q4 after subtracting the best-fit CO model \((N = 3.07 \times 10^{15} \text{ and } 3.49 \times 10^{15} \text{ cm}^{-2},\text{ respectively})\). The comet spectra were extracted from a region of 4″ width, centered on the comet nucleus. The red line is the predicted H\(_2\) fluorescence spectrum pumped by solar Ly\(\beta\) for a column density of \(1.0 \times 10^{14}\) cm\(^{-2}\) and a rotational temperature of 200 K. Other atomic contributions are shown in blue.

![Figure 8](image2.png)

Fig. 8.—Same as Fig. 5, for comet C/2001 Q4 (NEAT). The production rate for the native source model was \(1.76 \times 10^{28}\) molecules s\(^{-1}\).

![Figure 9](image3.png)

Fig. 9.—CO column density profile for comet C/2000 WM1 (LINEAR) extracted from regions of increasing widths, centered on the comet nucleus. The continuous line corresponds to the values predicted by the native source model with a production rate of \(3.56 \times 10^{26}\) molecules s\(^{-1}\), when integrated over similar slit sizes.
width. Fitting a native source model integrated over the same rectangular regions we obtain a CO production rate of \((3.56 \pm 0.2) \times 10^{26}\) molecules s\(^{-1}\). This model is shown by a solid line in Figure 9. A CO production rate of \(0.44\% \pm 0.03\%\) relative to water is obtained adopting the favored H\(_2\)O production rate for the FUSE observations \((8.0 \times 10^{28}\) molecules s\(^{-1}\); Weaver et al. 2002). This makes C/2000 WM\(_1\) (LINEAR) the most CO-poor comet of our sample.

4.2. HST GHRS Observations: C/1996 B2 (Hyakutake)

Since the HST GHRS observations do not provide spatial information, we are unable to derive a CO column density profile for comet C/1996 B2 (Hyakutake). The CO production rate is obtained by comparing the CO column density derived by fitting our fluorescence model to the GHRS spectrum with the CO column density predicted for the GHRS aperture by the native source model. As model parameters for the synthetic spectrum we used a rotational temperature of 72 K given by the \(63\tau_1\) dependence from DiSanti et al. (2003), which is similar to the value given by Lis et al. (1997), and a \(b\) parameter of 2.0 km s\(^{-1}\), within the range of outflow velocities measured by Wouterloot et al. (1998) but slightly higher than derived from the optical line widths of Combi et al. (1999). A lower \(b\) value is inconsistent with the total amount of absorbed solar radiation (from the conservation of the number of photons), suggesting larger turbulent motions in the \(1.74'' \times 1.74''\) area probed by GHRS.

Using the CO \(A-X\) fluorescence model with self-absorption we obtain a first estimate for the CO column density averaged over the GHRS slit. However, the predicted line ratios are not in agreement with the data (see the red line in Fig. 10, top). In order to improve the fit and better constrain the column density we adjust the model to account for the geometry of the Sun-comet-Earth system. Starting with the previous estimate on the CO column we can constrain the ratio between the line-of-sight column density and the absorbing column on the comet-Sun direction. Iterating this procedure we obtain a best-fit value for the line-of-sight CO column density of \((1.45 \pm 0.87) \times 10^{16}\) cm\(^{-2}\). This model is shown with a red line in the bottom panel of Figure 10. The blue model in the same figure contains contributions from atomic species C \(_{ii}\), C \(_{i}\), O \(_{i}\), and S \(_{i}\), as labeled, which account for the remaining features in the spectrum. Under the assumption of a native source model with the same lifetime and gas outflow velocity as employed for the STIS observations, we obtain a CO production rate of \((4.97 \pm 0.07) \times 10^{28}\) molecules s\(^{-1}\). This represents the highest CO production rate relative to water from our sample, 20.9\% \pm 0.3\%, using the water production rate of \(2.38 \times 10^{29}\) molecules s\(^{-1}\) from Combi et al. (1998).

Figure 10 also clearly shows the presence of several bands originating on the \(v' = 9\) and 14 levels that are pumped by solar O \(_{i}\) \(\lambda 1302\) and H \(_{i}\) Ly\(_{\alpha}\), respectively (Wolven & Feldman 1998). Kassel (1976) first pointed out that scattering of Ly\(_{\alpha}\) by the (14, 0) band is comparable to, if not larger than, the direct solar scattering by all other bands of the CO Fourth Positive system for CO column densities \(\geq 10^{17}\) cm\(^{-2}\). The (14, 4) band was subsequently identified in the spectrum of Venus (Durrance 1981). These bands are not detected in any of the STIS comet spectra. The high column density in the field of view of comet Hyakutake makes them visible, albeit at low signal-to-noise ratio, and allows for a determination of column density independent of optical saturation effects. In evaluating the fluorescence efficiency of these bands, the overlap between the solar lines and the individual lines of the CO bands is a sensitive function of rotational temperature and heliocentric velocity and requires accurate profiles of the solar lines (Lemaire et al. 2002). The (14–3), (14–4), and (9–2) bands are included in the model and are seen to be completely consistent with the CO column density derived above.
We note that the optically thick bands connecting to the $\nu' = 0$ level of the ground state, namely (1–0) at 1509.8 Å, (2–0) at 1477.6 Å, (3–0) at 1447.4 Å, (4–0) at 1419.1 Å, and (5–0) at 1392.6 Å, all labeled in Figure 10, do not show a significant variation due to geometric corrections. This can be understood from the fact that while the correction adjusts the line-of-sight column density relative to the column absorbing the solar radiation, the emission in the optically thick lines will still be determined by the column corresponding to one optical depth, which depends only on temperature and $b$ parameter. The (0–0) band at 1544.5 Å does not seem to follow this pattern, due to blending with the (3–2) band at 1542.5 Å. The same lack of variation is exhibited by the bands belonging to optically thin progressions pumped by solar emission lines, such as (14–3) at 1315.7 Å and (9–2) at 1377.8 Å. This is a result of the simple linear scaling in the optically thin limit, which makes the absorbing column indistinguishable from the emitting column.

5. DISCUSSION

5.1. Sources of Uncertainty and Comparison with Other Measurements

The derived column densities are subject to uncertainties due to the choice of model parameters and background subtraction. A more complete model would involve a multidimensional $\chi^2$ minimization to constrain simultaneously the column density, rotational temperature, and Doppler $b$ parameter. Aside from the fact that this approach requires rather large computational resources, we expect that given the quality of the data the resulting $\chi^2$ three-dimensional surfaces will have rather low contrast minima and the improvement in the resulting production rates would be negligible. More of a concern is the background subtraction in the HST STIS data. The background is variable both in the spatial and spectral directions from one observation to another, making it impossible to give a comprehensive subtraction prescription. The optimal background subtraction is determined on a case-by-case basis. While the background-related uncertainties can lead to $\sim$30% variations in the values for the column densities, we estimate that the change in the resulting production rates is only about 6%. These values are only slightly larger than the 1 $\sigma$ error bars from the $\chi^2$ minimization. Similarly, increasing the rotational temperature from 70 to 100 K results in a $\sim$30% decrease in production rate. However, the rotational temperatures relevant for our observations were derived from either FUSE observations or from the $T_{\text{rot}}$ versus $r_h$ dependences obtained by radio measurements, and are constrained to better than $\pm 6$ K. This results only in a $\sim$7% variation in production rate, as the heliocentric distance of the comet does not vary significantly during our observations. The absolute values of the CO production rate and column density are also sensitive to the STIS calibration pipeline, which is based on point-source stellar standards. The column densities derived from STIS data could be overpredicted by at most 30% due to calibration offsets. Variations in the parameters used for the native source model, such as the CO lifetime and outflow velocity, could also change the production rate by a few percent. Other less quantifiable uncertainties to which the fluorescence model is particularly sensitive are the oscillator strengths and the UV solar flux, especially due to the variable emission features at high solar activity. Overall, we estimate that the systematic errors in the production rates amount to at most 15%.

To assess the effects of the error sources mentioned above, it is useful to compare our results to other measurements of the CO production rate in these comets from different spectral regions. The values obtained for comets C/2000 WM$_1$ (LINEAR) and C/1996 B2 (Hyakutake) are in excellent agreement with previous measurements. For comet C/1996 B2 (Hyakutake) at similar heliocentric distances, Biver et al. (1999) find values of $(4.98 \pm 0.09) \times 10^{28}$ molecules s$^{-1}$ (0.952 AU) and $(4.84 \pm 0.58) \times 10^{28}$ molecules s$^{-1}$ (0.894 AU). The production rate relative to water is comparable to previously measured mixing ratios of 14%–19% (DiSanti et al. 2003) and 22% (Biver et al. 1999). For comet C/2000 WM$_1$ (LINEAR), both FUSE (Weaver et al. 2002) and radio (Biver et al. 2006) observations are consistent with a CO mixing ratio of $\sim$0.4% and a CO production rate of $\sim(3–4) \times 10^{28}$ molecules s$^{-1}$. A good agreement is again found when comparing the CO production rate derived for comet C/2001 Q4 (NEAT) with the results from FUSE observations (§ 4.1.2), listed in Table 3. We note that the measurements based on FUSE observations used only the cold component of CO (see § 4.1.2) in estimating the CO production rate. The cold component is believed to reflect the native source of CO, which is directly probed by STIS. The two values agree marginally within the error bars, and the mismatch can be attributed to short time variability and pointing instability for the FUSE observations. The count rates for the $C - X$ band in the FUSE data show factor of 2 variations with a periodicity of $\sim$20 hr. The FUSE observation overlaps with a much shorter STIS exposure, but the exact correlation in comet activity between the two data sets is hard to assess due to the different fields of view of the two instruments.

For comet 153P/Ikeya-Zhang the derived CO production rate is about a factor of 2 higher than the estimate from the range of values given by Biver et al. (2006; using the $r_h^{2.1}$ scaling, see Table 3) and DiSanti et al. (2002). The factor of 2 difference may be attributed to the uncertainties in the model parameters (rotational temperature and Doppler parameter) and in the background subtraction. However, as discussed above, the combined sources of error should result in less than 15% uncertainty. On the other hand, the $r_h^{2.1}$ dependence was derived from two measurements, one at 0.51 AU and the other at 1.26 AU, not excluding the possibility of temporal variations affecting our observations at 0.89 AU. Moreover, the uncertainties in the $r_h^{2.1}$ scaling have not been quantified, making our value for the production rate more reliable. The CO abundance relative to water is again higher than other estimates. However, using the water production rate given by the $19.0 \times 10^{28} \times r_h^4$ molecules s$^{-1}$ dependence derived by Biver et al. (2006) from H$_2$O observations made by the Odin satellite (Lecacheux et al. 2003), the CO mixing ratio decreases from 7.2% $\pm$ 0.4% to 5.1% $\pm$ 0.3%. This would place the CO relative abundance in comet 153P/Ikeya-Zhang closer to the value of 3.8% $\pm$ 0.6% at $\sim$1 AU (Biver et al. 2006) and 4.7% $\pm$ 0.8% at 0.78 AU (DiSanti et al. 2002).

5.2. Application to Broadband Imaging

The advantage of long-slit spectroscopy with STIS resides in giving us a better understanding of the spatial distribution of CO and the effects of increasing optical depths on the observed line ratios. We show how this information is applicable to broadband imaging by integrating the brightness of the $A - X$ bands in the STIS bandpass that are not contaminated by atomic lines. We then integrate the fluorescence model over the same bands, deriving an equivalent $g$-factor (as a ratio between the total brightness and the column density). Given the specific optical depth corrections in our model, the resulting $g$-factor will be in fact a brightness–column density dependence rather than a globally constant ratio. The radial brightness profile of the integrated bands observed by STIS can then be converted into a column density profile. This procedure can be visualized in Figure 11, where the data for comet 153P/Ikeya-Zhang have been used. The brightness...
thought to be related to the local gas and dust composition and temperature where they formed. Under the assumption that CO has been trapped by water ice during comet formation, the local temperature can be estimated from the observed production rate (Notesco et al. 1997). The formation temperatures in our sample are expected to range from 50 K or less, for comets C/2001 Q4, 153P/Ikeya-Zhang, and Hyakutake, to more than 60 K for C/2000 WM₃. If the gas trapping mechanism is more efficient than deposition, as suggested by laboratory studies of CO and water interactions (Collings et al. 2003), important enrichments for the species with a higher sublimation temperature can occur (Notesco et al. 1997). This can explain the lack of correlation between the abundance of CO and that of other molecules with a similar sublimation temperature (Gibb et al. 2003; Biver et al. 2002).

6. SUMMARY

We have developed a fluorescence model for the interpretation of CO \( A \rightarrow X \) emission observed in several recent comets by the Hubble Space Telescope employing the latest values for the transition wavelengths and oscillator strengths. The radiative transfer approximation takes into account saturation effects using Voigt profiles for each of the \( \sim 10^3 \) transitions. Self-absorption is introduced using a photon mean free path approximation. This process is significant for column densities above a few \( \times 10^{14} \) cm\(^{-2} \), encountered close to the comet nucleus. It is shown that the model reproduces the optically thin limit for lower column densities, at larger distances from the comet center. When the column densities are of the order \( \sim 10^{16} \) cm\(^{-2} \) or above, a good fit to the data requires a distinction between the absorbing CO column in the Sun-comet direction and the emitting column along the line of sight. The approximations in the radiative transfer model are justified for the quality of the available data. Constraining better the optical depth effects demands an exact treatment, which would be warranted at higher spectral resolution in order to predict the relative strengths of individual lines contained in the \( A \rightarrow X \) bands.

For the comets observed by STIS the fit of the column density profile is consistent with a predominantly native source, with production rates ranging from \( 3.56 \times 10^{26} \) to \( 1.76 \times 10^{24} \) molecules s\(^{-1} \). The quality and the spatial extent of the STIS data does not allow for a detection of a small extended source component. We find large variations in the CO abundance relative to water, from 0.4% for C/2000 WM₃ (LINEAR) to 21% for C/1996 B2 (Hyakutake). This diversity among comets is still not well understood, as no clear trend emerges at the current stage of comet observations (Bockelée-Morvan et al. 2004). In spite of the caveats discussed in § 5, the absolute values of the CO production rate and column density are better constrained by the introduction of optical depth effects and give reasonable confidence levels for the CO production rate and a good agreement with previous results. Moreover, the present model proves to be a valuable tool in analyzing broadband imaging data, such as the SBC ACS observations of comet 9P/Tempel 1 (Feldman et al. 2006a).

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