Soft X-ray Images of the Solar Corona with a Normal-Incidence Cassegrain Multilayer Telescope

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High-resolution images of the sun in the soft x-ray to extreme ultraviolet (EUV) regime have been obtained with normal-incidence Cassegrain multilayer telescopes operated from a sounding rocket in space. The inherent energy-selective property of multilayer-coated optics allowed distinct groups of emission lines to be isolated in the solar corona and the transition region. The Cassegrain telescopes provided images in bands centered at 173 and 256 angstroms. The bandpass centered at 173 angstroms is dominated by emission from the ions Fe IX and Fe X. This emission is from coronal plasma in the temperature range $0.8 \times 10^6$ to $1.4 \times 10^6$ K. The images have angular resolution of about 1.0 to 1.5 arc seconds, and show no degradation because of x-ray scattering. Many features of coronal structure, including magnetically confined loops of hot plasma, coronal plumes, polar coronal holes, faint structures on the size scale of supergranulation and smaller, and features due to overlying cool prominences are visible in the images. The density structure of polar plumes, which are thought to contribute to the solar wind, has been derived from the observations out to 1.7 solar radii.

Research Articles

T he study of the Solar atmosphere is complicated in that the sizes of the fundamental structures that control the important physical processes are below the resolution limit possible with instruments now available for all wavelengths. This problem is a consequence generally of the dominant influence of the solar magnetic field, which is characterized by structure on the scale of 70 kilometers (0.1 arc second as viewed from Earth) or less (1). The problem has been especially acute for soft x-ray observations because the available techniques could not provide observations with both high spatial and high spectral resolution simultaneously.

Most solar x-ray observations have been made with grazing-incidence optical systems (2). These systems, which were first developed 40 years ago (3, 4), have been used successfully to observe the solar corona (5) and to study cosmic x-ray sources (6). Grazing-incidence optics can provide high-quality images, but they have no inherent ability to resolve spectral features. Consequently, images obtained with grazing-incidence optics and thin filters are of low to moderate spectral resolution, and they cannot be used to isolate...
material at a single temperature or to resolve macroscopic dynamical effects (which require the study of line profiles). Observations with dispersive spectrometers (7, 8) can provide high spectral resolution, but resolution of the physical structures responsible for the emission is limited.

Significant advances in the development of normal-incidence x-ray optics (9) now allow their use for solar observations. Normal-incidence x-ray telescopes have mirrors coated with multiple thin atomic layers ("multilayers"). These layered structures act as Bragg diffractions for soft x-ray and extreme ultraviolet (EUV) radiation (1 to 1000 Å). Multilayer structures, when deposited on substrates of sufficient smoothness and optical precision, can act as optical elements in scanning microscopes (10), spectroscopic instruments (11), and astronomical telescopes (12). Multilayer optics have already had a significant impact on laboratory studies of solids (13), biological systems (14), and plasmas (15).

We have used normal-incidence multilayer optics in a Cassegrain telescope to obtain images of the solar corona in the 171 to 175 Å bandpass. These high-resolution images were obtained from a Nike-Black Brant sounding rocket launched at 1805 universal time (UT) on 23 October 1987 from the White Sands Missile Range, New Mexico. Many coronal structures are visible in the images (Fig. 1, A and D). In this article, we compare these images with images obtained at other wavelengths to indicate how the structure of the solar plasma varies with temperature (Fig. 1, A, B, and C) and how the structure is shaped by solar magnetic fields (Fig. 1, E and F). The Cassegrain telescope images demonstrate that multilayer optics can furnish high-resolution images and spectroscopic analysis simultaneously, a capability that has only now become available to x-ray astronomers and solar physicists (16).

**Multilayer x-ray—EUV optics.** Grazing-incidence high-resolution x-ray telescopes and microscopes have been designed almost

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**Fig. 1.** (A) The solar corona as photographed in the emission of the resonance lines of Fe IX at 171 Å and Fe X at 174.5 Å, on 23 October 1987 at 1809 UT. The length of the exposure was 64 seconds. The dark bands visible at northern latitudes are due to overlying cool prominences, which are also clearly visible in (B). The image corresponds to coronal structures from approximately 0.8 × 10⁴ to 1.4 × 10⁵ K. (B) The sun as photographed in the light of the Balmer-α (Hα) absorption line of hydrogen at 6563 Å with a bandpass of 0.5 Å at 1510 UT on 23 October 1987. The core of the Hα line is formed in the chromosphere in the temperature range 6000 to 7000 K. [Photograph courtesy of D. Speich, National Oceanic and Atmospheric Administration, Boulder, Colorado] (E) The solar chromosphere at 1579 UT on 23 October in the K3 absorption line of Ca II (λ = 3934 Å). The photograph was made with a 0.3 Å wide filter. The Ca II K3 line, which is formed at approximately 1 × 10⁵ K, displays the structure of the chromospheric network (46). [Photograph courtesy of B. Labonte and D. Braun, University of Hawaii Mees Solar Observatory, Haleakala, Maui, Hawaii] (D) Locations of various phenomena in A. The fiducial marks are at 1-arc-minute intervals. (E) Kitt Peak magnetogram for 23 October (1725 UT) showing the fine structure of the solar magnetic field. The gray indicates neutral regions, black negative polarity, and white positive polarity. The active regions, seen in the 173 Å image where emission from the magnetically confined coronal plasma is most intense, are clearly associated with bipolar regions with strong magnetic fields. The resolution of the magnetogram is 1.0 arc second. [Photograph courtesy of Jack Harvey of the National Solar Observatory, Tucson, Arizona] (F) Stanford magnetogram for 24 October 1987 (2124 UT) showing the large-scale structure of the solar magnetic field. The dark lines indicate the location of the neutral lines. The solid contours indicate positive polarity, and the dashed contours indicate negative polarity. The resolution of the magnetogram, which is intended to portray large-scale magnetic structure, is 3 arc minutes. Because of inclement weather, no magnetogram was obtained on 23 October. [Photograph courtesy of Phil Scherrer, Stanford University Wilcox Solar Observatory, Stanford, California]
exclusively with deep conic mirrors (4). When soft x-rays strike a mirror at a grazing angle that is smaller than the critical angle (typically 0.5 to 3 degrees), they are efficiently reflected. The mirrors in the typical grazing-incidence telescope (the Wolter type I telescope (17)) present only a thin annulus for collecting incoming radiation. As a result, mirrors of large diameter and with long focal lengths are required to achieve reasonable collecting areas. Some improvement can be achieved by nesting one mirror inside another. The point response functions of grazing-incidence mirrors have broad wings (18), which make resolving small, faint structures difficult. As a result, radiation that is nonspecularly scattered by mirror surface finish imperfections and particulate and condensible contaminants also commonly degrades the performance of grazing-incidence optics. Resolution below an arc second may be difficult to obtain with optical systems composed of grazing-incidence optics alone.

The use of normal-incidence optics allows many of the above problems to be eliminated. For a given aperture the collecting area and field of view of normal-incidence telescopes are significantly larger than those for grazing-incidence telescopes. Compound normal-incidence optical systems can be folded and thus are more compact. Normal-incidence optical surfaces are significantly easier to manufacture than grazing-incidence ones. The aberrations resulting from violations of the Abbé sine condition that are associated with grazing-incidence optical (18) and spectroscopic (19) systems can be avoided. In addition, because normal-incidence optics do not rely on grazing reflection angles to focus radiation, image degradation resulting from nonspecular scattering is much less severe.

Normal-incidence x-ray telescopes are made by coating (in vacuo) ultrasmooth surfaces with a series of alternating layers of two different materials, one a strong x-ray scatterer with high atomic number, and the second a weak x-ray scatterer or "spacer" with low atomic number. This synthetic structure acts as an x-ray Bragg diffractor. X-ray reflectivity is maximized when the condition for constructive interference is satisfied as indicated by the Bragg equation, \(n \lambda = 2d \sin \theta\), where \(n\) is the order of diffraction, \(\theta\) is the angle of incidence with respect to the surface, \(d \) is the effective layer-pair thickness (corrected for refraction effects), and \(\lambda\) is the wavelength of the incident radiation. Although the reflectivity at each interface is small, when tens to hundreds of layer pairs are used, excellent reflectivities (more than 50 percent near normal incidence) can be achieved. Methods for fabricating these thin-film synthetic microstructures (multilayers) were developed by Spiller (20) and by Barbee (21).

In contrast to naturally occurring Bragg crystals, multilayer structures can be deposited on curved substrates of sufficient surface smoothness, which are required for high-performance optical systems. With proper selection of the materials and control of the thicknesses of the scatterer and spacer layers, multilayer structures can also be fabricated to have any layer-pair spacing between 15 and 500 Å. Hence, conventional astronomical telescope configurations (22) with normal-incidence reflection optics coated with a multilayer of the appropriate layer-pair spacing can be made to operate in the soft x-ray-EUV wavelengths where in the past only grazing-incidence or single-reflection normal-incidence optical systems could be used.

Because multilayer surfaces are reflective only in a certain bandpass, they can be used to isolate a particular region of the spectrum or a particular emission line or line multiplet between 30 and 1000 Å (at normal incidence). Their effectiveness is limited at short wavelengths by substrate smoothness and the atomic nature of matter and at long wavelengths by absorption in the multilayer structure. Compound systems containing both grazing-incidence and multilayer optics at non-normal incidence are required in order to operate at wavelengths less than 30 Å (12). The spectral resolution of a multilayer is determined by the number of layers that participate in the Bragg diffraction process, which in turn depends on the total number of layers present and the amount of attenuation in each layer. For multilayer-coated optics, the spectral bandpass \((\Delta \lambda / \lambda)\) can be as large as 20 percent or as small as 1 percent. Normal-incidence reflection efficiencies can be made as high as 70 percent with the proper choice of materials.

The first successful imaging experiments with normal-incidence multilayer optics were performed in the laboratory during the early 1980's (23). Subsequently, several groups of investigators have successfully fabricated multilayer-coated optics for the soft x-ray-EUV (24). Underwood et al. (25) have obtained a moderate resolution (about 10 arc seconds) image of a single solar active region with a single-reflection (Hercchelian) multilayer telescope having a bandpass centered at 44 Å (corresponding to Si XII radiation).

**Instrumentation.** We obtained our normal-incidence images of the sun, using a Cassegrain telescope mounted on the Stanford-Marshall Space Flight Center (MSFC) rocket spectroheliograph instrument. This instrument contained nine optical systems: four systems, including the Cassegrain telescope, had normal-incidence optics exclusively, four had normal-incidence optics in conjunction with grazing-incidence optics, and one had grazing-incidence optics exclusively (12, 26, 27). The Cassegrain telescope had a bandpass centered at 173 Å. The primary mirror was a concave spherical mirror with a diameter of 6.4 cm and a radius of curvature of 1.2 m. The secondary mirror was a 2.5-cm-diameter convex spherical mirror with a radius of curvature of 0.5 m. This configuration yielded a focal length of 2.0 m for the combined f-33 system. We used spherical primary and secondary mirrors instead of conic mirrors because of the availability of spherical optics with superb surface quality; the resolution limit of the spherical mirrors [0.5 arc second root mean square (rms)] was comparable to that of the film used. A mesh-supported, 1720 Å thick aluminum filter (28) excluded visible and ultraviolet light from the camera and also helped define the bandpass.

The wavelength sensitivity \((e)\) of the Cassegrain telescope was

![Fig. 2. Spectral response \((e)\) of the multilayer Cassegrain telescope, superimposed on the solar spectrum (30).](image-url)
A MSFC collimated beams contribute less IX Cassegrain with was chosen of each ability bandpass. Improved the measured visible light resolution at 17 arc minutes off-axis is due to diffraction effects caused by vignetting.

The images were recorded on an EUV-sensitive, tabular grain (T-grain) (39) emulsion that was specially prepared for us by the Eastman Kodak Company. This emulsion is similar to Kodak TMAX 100, but was prepared without a gelatin overcoat for optimum performance in the EUV. The 35-mm film had a Rem-Jet backing which prevented static discharges during film transport in the vacuum; the film was transported in conventional Canon T-70 cameras with no adverse effects. Previous observations in the EUV region have been carried out with Schumann emulsions (40), which are extremely sensitive to contaminants and abrasions and yield a maximum photographic diffuse density no greater than 1.6. Our T-grain emulsion yielded densities greater than 3.0 (semispecular densities greater than 3.9), and the T-grain 100 film provided high sensitivity throughout the spectral region studied during this flight (8 to 256 Å). The film obtained from the 171 to 175 Å Cassegrain was processed to yield a visible light ASA (American Standards Association) speed equivalent of 400, thus providing optimum sensitivity to faint features in the far corona. Hoover et al. (41) have discussed the advantages of these new T-grain films at soft x-ray–EUV wavelengths.

Observations. A wide range of coronal features are visible in the bandpass between 171 and 175 Å (Figs. 1A and 5). These images complement the information that can be obtained (i) from images in the spectral region from 1 to 30 Å (7), which is most effectively studied with grazing-incidence optical systems (42), and (ii) from the spectral region between 304 and about 600 Å, (43) which is accessible to conventional single-reflection optics (16, 44). Below 30 Å, emission of the nonflaring solar atmosphere is dominated by hot (≥ 2.0 × 10^6 K), dense (about 10^14 ions per cubic centimeter) plasma confined in magnetic flux tubes (45). Between 304 Å and about 600 Å, solar emission is dominated by plasma from 2 × 10^5 K to 1 × 10^6 K (the transition region (46, 47)), which is strongly confined by the highly structured magnetic field of the chromospheric network (Fig. 1C).

Changes in the structure of the atmosphere occur between 1 × 10^6 K and 1.5 × 10^6 K. In closed-field regions much of the gas at these temperatures is confined in loop structures (45). Systems of coronal loops are visible on the limb in the northwest, west, and east (Figs. 1, A and D, and 5), and are responsible for the emission visible from active regions 4872 and 4873 (see also Fig. 1E), at south 60°, west 0° to 50°. In open-field regions, where the coronal plasma is expanding into the solar wind, structures such as streamers and polar plumes (47) that are at the interface between the corona and the interplanetary medium, are visible in the 171 to 175 Å images at 1 × 10^6 K to 1.5 × 10^6 K (Figs. 1, A, and D, and 5). Streamers, representing the corona–solar wind interface, are visible virtually everywhere on the limb, including coronal plumes in the polar regions, which are otherwise deficient in emission. Coronal holes, representing regions that generally lack closed magnetic structures capable of confining hot plasma, are clearly visible at the poles.

Other features that are visible in the 171 to 175 Å images are regions of low emission that coincide with large-scale loops of cool gas embedded in the hotter coronal gas that produces the 171 to 175 Å emission. These loops of gas at chromospheric temperatures (6 × 10^5 K to 3 × 10^6 K typically) are called prominences (48), and can be observed in emission above the limb (Fig. 1, A and D). Prominences observed in absorption on the disk (for example, in the Balmer-α (Hα) line images at 6563 Å (Fig. 1B)) are called filaments. Spectacular examples of such structures have been ob-
served in He II Lyman-α line images at 304 Å (16). Some of the 
"prominence cavities" (or filaments) seen in our images are also 
clearly visible in the Hα chromospheric image (Fig. 1B). A dark 
feature that crosses the northwest limb at 70° latitude (Fig. 1A) 
can be seen in emission above the limb in the Hα image (Fig. 1B). 
Because they are cooler, prominences represent material that is 
considerably denser than the surrounding coronal gas. The cool gas 
is supported against gravity by the magnetic field, and hence 
prominences occur at locations where the field is horizontal (that is, 
where the vertical component of the field is zero). These "neutral 
lines," which are closely delineated on the large-scale magnetic field 
map (Fig. 1F), are well correlated with regions of low emission 
in the 171 to 175 Å images. Although the depressed emission at 171 to 
175 Å associated with prominence cavities may be due in part to the 
exclusion of hot gas from the vicinity of the neutral line, the abrupt 
decrease in emission at the northwest limb associated with the 
prominence at 70° latitude suggests that prominence gas is also 
abrupting 171 to 175 Å emission from hot material lying below or 
behind the prominence loop. Considerably more prominence structu 
re is visible in the 171 to 175 Å images than in the Hα image, 
possibly because some of the prominence material is at temperatures 
where all of the hydrogen is ionized, and unable to absorb Hα 
radiation. The absorption we observe is occurring in the He II 
continuum, shortward of 227 Å. Schmahl et al. and Schmahl and 
Orrall (49) have commented on a similar phenomenon occurring in 
the H I continuum. The large absorption feature in the southeast 
of our images at 45° latitude on the limb is an example of this 
phenomenon, and may represent a giant helmer streamer (50).

Another feature of interest is the structure extending toward the 
northeast from a point on the southeast limb at 40° latitude. This 
structure could represent a coronal spray, material ejected from 
the region of a flare behind the limb in the low corona. (There is no 
record of a flare occurring on the visible side of the limb near the 
time of our observation.) Such structures have been observed in He 
II Lyman-α images (16); our observations indicate that large sprays 
in radiation may also be observed from an ion present in plasma at 1 
× 10⁶ K.

The images of Figs. 1A and 5 reveal that the corona is very 
complex at 1 × 10⁶ K, with cool material embedded in the hot 
corona. Among the structures more clearly resolved in our images 
compared to earlier images are coronal loops on various size scales, 
polar plumes in emission, and prominence cavities on the disk and 
above the limb. Fine structure that appears to be related to 
supergranulation can also be seen on the disk (Fig. 5). Further 
analyses should yield information on a number of critical problems 
off coronal structure, including the nature of the chromospheric 
network in the lower corona, the structure of the corona—solar wind 
interface, the nature of the prominence-corona interface, and the 
nature of coronal loops.

**Polar plumes.** Polar plumes are believed to be the major source of the 
high-speed solar wind streams associated with coronal holes. 
Withbroe (51) has reviewed models of polar plumes and pointed out 
the importance of high-resolution (~1 arc second) observations of 
polar plumes to an understanding of the physics of coronal holes. 
Previous observations of polar plumes (on Skylab in the soft x-ray 
and EUV (52–54), have been limited in resolution to about 3 to 5

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**Table 1. Solar lines within the Cassegrain bandpass**

| Wavelength (Å) | Intensity (I) (10⁻⁴ ergs cm⁻² sec⁻¹) | Temperature (10⁶ K) | Ion | Contribution (percentage) |
|----------------|-------------------------------------|-------------------|-----|---------------------------|
| 167.495        | 2.5                                 | 0.80              | Fe VIII | 0.10                     |
| 168.176        | 4.6                                 | 0.80              | Fe VIII | 0.20                     |
| 168.548        | 2.0                                 | 0.80              | Fe VIII | 0.10                     |
| 168.933        | 1.6                                 | 0.80              | Fe VIII | 0.07                     |
| 171.075        | 51.5                                | 1.00              | Fe IX   | 45.9                     |
| 171.359        | 5.0                                 | 2.50              | Ni XIV  | 4.8                      |
| 172.174        | 1.1                                 | 0.25              | O V     | 1.1                      |
| 172.936        | 1.4                                 | 0.30              | O VI    | 1.3                      |
| 173.081        | 3.0                                 | 0.30              | O VI    | 2.9                      |
| 173.354        | 52.7                                | 1.25              | Fe X    | 37.3                     |
| 175.266        | 4.5                                 | 1.25              | Fe X    | 2.1                      |
| 175.474        | 1.3                                 | 1.25              | Fe X    | 0.45                     |
| 176.694        | 2.5                                 | 2.80              | Ni XV   | 0.15                     |
| 177.243        | 29.0                                | 1.25              | Fe X    | 0.70                     |
| 178.060        | 2.0                                 | 1.55              | Fe XI   | 0.04                     |
| 179.762        | 1.0                                 | 1.55              | Fe XI   | 0.05                     |
| 180.407        | 59.0                                | 1.55              | Fe XI   | 2.50                     |
| 180.600        | 2.0                                 | 1.55              | Fe XI   | 0.07                     |
| 181.140        | 2.7                                 | 1.55              | Fe XI   | 0.07                     |
| 186.885        | 9.8                                 | 2.00              | Fe XII  | 0.08                     |
| 188.219        | 39.8                                | 1.55              | Fe XI   | 0.22                     |

*From (36). †Measured by Malinovsky and Heroux (37) except for λ = 171.359 and 179.762 Å, $e_i^x$, $e_j^y$, where i and j are solar lines and the sum is over all the lines in the table; $e$ is taken from Fig. 4. $§$Estimated from the computations of Doschek and Cowan (36).
arc seconds, and to distances of 0.4 solar radii or less above the limb. To demonstrate the power of multilayer optics for the study of the structure of the interface between the corona and the solar wind, we examined the structure of a typical polar plume (Fig. 5). We calculated plume densities assuming that the plume is isothermal with a temperature of about $1.1 \times 10^6$ K, a value consistent with previous analysis of Skylab observations. The mean electron density $\langle n_e \rangle$ in the plume in a resolution element $dx dy$ can then be derived from the coronal excitation equation (7), which expresses the intensity in any emission line $\lambda_j$ as

$$I_{\lambda_j} = \frac{\hbar c}{\lambda_j} a_{\lambda_j} A_2 \int dx \int dy \int ds a_Z(T_e) a_{\lambda_j}(T_e) n_e^2$$

$$= \frac{\hbar c}{\lambda_j} \langle n_e \rangle a_{\lambda_j} A_2 \int dx \int dy \int ds a_Z(T_e) a_{\lambda_j}(T_e)$$

where $s$ is the coordinate along the line of sight. This equation is based on the assumption that the observed emission is due entirely to collisional excitation (given by the temperature-averaged excitation function $a_{\lambda_j}$) from the ground level $j$ to the excited level $i$, and that the ionization temperature is equal to the electron temperature. The quantities $A_Z$ (abundance of element $Z$), $a_{\lambda_j}$ (population of ionization stage $z$ of element $Z$), $a_{\lambda_j}$ (number of hydrogen atoms per electron), and the excitation functions are taken from Meere and Gronenschild (55). Only the strongest lines in the bandpass, Fe IX (171 Å), Fe X (174.5 Å), O V (172.2 Å), O VI (173 Å), and Ne IV (172.6 Å), were included in the computations. The derived (Fig. 6) are in good agreement with earlier EUV and soft x-ray results (52-54) and with earlier eclipse results (56) for less than 1.4 solar radii, and extend the EUV analysis of plume structure to about 1.7 solar radii (57). The plumes near the edge of the polar coronal holes tilt outward toward lower latitudes as compared to the radial direction. The plume that is plotted in Fig. 6 is tilted about 20° away from the radial direction.

Future applications of normal-incidence x-ray optics. Our data demonstrate that normal-incidence multilayer x-ray telescopes can form high-resolution images with significantly lowered levels of x-ray scatter. Furthermore, these systems are not subject to acute optical aberration effects over reasonable fields of view. We anticipate that multilayer telescopes will provide solar x-ray images with spatial resolution approaching 0.1 arc second in the near future in well-defined spectral bands with excellent sensitivity. Additional capabilities that can be anticipated with multilayer optics, such as multilayer gratings, zone plates, and compound systems incorporating grazing-incidence primary mirrors and multilayer magnifying optics (27) or multilayer Rowland circle grating or “crystal” spectrometers (19) should extend the utility of multilayer optics to shorter wavelengths and provide astronomers with powerful high spectral resolution instruments that are stigmatic. These instruments should allow the study of fundamental problems in solar physics such as (i) the structure, dynamics, and energy balance (including the heating source or sources) of coronal loops, (ii) the mechanism responsible for generating the solar wind, and (iii) the configurations in the corona which are necessary for the onset of flares.

The application of multilayer optics to nonsolar x-ray astronomy should prove extremely fruitful, as well. X-ray astronomy has revealed the existence of high-temperature astrophysical phenomena that range from thousands of kilometers to megaparsecs (58-60). Although not as bright from Earth as the sun, many cosmic x-ray sources have complex extended thermal and spatial structure. Among such sources are supernova remnants (61), the hot component of the interstellar medium in our galaxy (62) and in other nearby galaxies, and the hot intracluster gas bound in clusters of galaxies (60). The study of the thermal and density structure of these extended cosmic plasmas may be possible with the use of multilayer optics coupled to grazing-incidence optical primary mirrors such as the Advanced X-Ray Astrophysics Facility (63). In addition, spectrally resolved images of the halos (64) caused by interstellar scattering, which surround the image of every compact x-ray source, can be used to address fundamental questions concerning the composition and structure of the interstellar dust (65).

In summary, we have used normal-incidence multilayer optics to produce high spatial resolution full disk images of the sun in selected narrow wavebands in the soft x-ray-EUV. The photographic densities observed even in short exposures clearly reveal the high reflectivity and excellent waveband match of the primary and secondary multilayer mirrors in the Cassegrain telescope and also demonstrate the superiority of the new Eastman Kodak T-Grain emulsions over Schumann emulsions, which have long served as the photographic film of choice for soft x-ray-EUV observations (66). Multilayer optics can also lessen x-ray telescope system design requirements and yield improved image quality in an experimentally difficult spectral region. Multilayer optics do not exhibit the x-ray scattering characteristic of grazing-incidence telescopes. Consequently, small faint features can be clearly seen even though they are near bright loops or active regions. Because of these important technological developments, the images obtained during this flight have revealed features of the solar corona in greater detail at soft x-ray-EUV wavelengths. Clearly, multilayer optics will play a profound role in laboratory and astronomical x-ray imaging.

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