FUSE spectroscopy of sdOB primary of the post common-envelope binary LB 3459 (AA Dor)*

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ABSTRACT

Context. LB 3459 (AA Dor) is an eclipsing, close, post common-envelope binary (PCEB) consisting of an sdOB primary star and an unseen secondary with an extraordinarily low mass (M2 ∼ 0.066 M⊙) – formally a brown dwarf. A recent NLTE spectral analysis shows a discrepancy with the surface gravity, which is derived from analyses of radial-velocity and lightcurves.

Aims. We aim at precisely determining of the photospheric parameters of the primary, especially of the surface gravity, and searching for weak metal lines in the far UV.

Methods. We performed a detailed spectral analysis of the far-UV spectrum of LB 3459 obtained with FUSE by means of state-of-the-art NLTE model-atmosphere techniques.

Results. A strong contamination of the far-UV spectrum of LB 3459 by interstellar line absorption hampers a precise determination of the photospheric properties of its primary star. Its effective temperature (Teff = 42 kK) was confirmed by the evaluation of new ionization equilibria. For the first time, phosphorus and sulfur have been identified in the spectrum of LB 3459. Their photospheric abundances are solar and 0.01 times solar, respectively. From the C IV λ1174 – 1177 Å multiplet, we can measure the rotational velocity vrot = 35 ± 5 km/sec of the primary of LB 3459 and confirm that the rotation is bound. From a re-analysis of optical and UV spectra (analogue to Rauch 2000), we determine a slightly higher surface gravity log g = 5.3 ± 0.1 compared to Rauch (2000, log g = 5.2 ± 0.1).

Conclusions. The rotational velocity of the primary of LB 3459 is consistent with a bound rotation. The higher log g reduces the discrepancy in mass determination in comparison to analyses of radial-velocity and lightcurves. However, the problem is not completely solved.

Key words. Stars: abundances – Stars: atmospheres – Stars: binaries: eclipsing – Stars: early-type – Stars: low-mass, brown dwarfs – Stars: individual: AA Dor, LB 3459

1. Introduction

The eclipsing binary system LB 3459 (AA Dor) is a blue foreground object of the LMC at a spectroscopic distance of d = 396 pc (Rauch 2000). It consists of a sdOB with an effective temperature of Teff = 42 kK and a low-mass companion from which no direct spectroscopic information has been obtained yet. Due to its mass of about 0.066 M⊙ it formally lies in the brown-dwarf range. The orbital period is about 0.26 d and the inclination is i = 90°.

The system was analyzed several times, starting with Kilkenny et al. (1979, 1981) and Paczynski (1980), who established the basic parameters. First NLTE analyses of the primary were performed by Kudritzki et al. (1982) and Lynas-Gray et al. (1984). Details on the history of this object can be found in Rauch (2000, 2004).

Recent spectral analyses by means of NLTE model-atmospheres that were based on optical (ESO CASPEC1) and ultraviolet (IUE2) observations (Rauch 2000) have shown a discrepancy with analyses of radial-velocity and lightcurves (Hilditch et al. 1996, 2003). Rauch (2000) determined a surface gravity of log g = 5.21 ± 0.1, while Hilditch et al. (2003) obtained log g = 5.45 – 5.51 from lightcurve and mass function. Because the analysis of Rauch (2000) suffered from the long exposure times (some hours) hence smearing, due to the orbital movement of the available spectra, Rauch & Werner (2003) measured the radial-velocity curve from optical spectra with short (180 sec) exposure times obtained with ESO’s VLT4 and UVES. They obtained a radius r1 = 0.169 R⊙ which is smaller than the r1 = 0.236 R⊙ found by Rauch (2000). Since the stellar radius depends on g this may be a hint of a higher value than log g = 5.21. However, data reduction of the Balmer lines in the UVES spectra used by Rauch & Werner (2003) was not very accurate, so a precise determination of log g of the primary by means of NLTE modeling techniques is still lacking.

Consequently, high-resolution and high-S/N observations (exposure times of 200 sec each) in the far-UV range were performed with the FUSE satellite. The FUSE wavelength range (904 Å < λ < 1187 Å) covers the hydrogen Lyman series except for Ly α. The series decrement is a sensitive indicator for log g.

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1 European Southern Observatory, Cassegrain Echelle Spectrograph
2 International Ultraviolet Explorer
3 Very Large Telescope
4 Ultraviolet and Visual Echelle Spectrograph
5 Far Ultraviolet Spectroscopic Explorer
A detailed spectral analysis of the FUSE observation of LB 3459 is described in Sect. 2. Since the far-UV spectrum of LB 3459 turned out to be strongly contaminated by interstellar absorption, we modeled the ISM line absorption (Sect. 4) in order to distinguish weak lines of iron-group elements (here: Ca, Sc, Ti, V, Cr, Mn, Fe, Co, Ni; Sect. 4).

2. The far-UV spectrum of LB 3459

The FUSE instrument consists of four independent, co-aligned telescopes and spectrographs. Taken together, the four channels span the wavelength range 904 Å < λ < 1187 Å with a typical resolving power of R ≈ 20 000. Further information on the FUSE mission and instrument can be found in Moos et al. (2000) and Sahnow et al. (2000).

Far-UV observations were performed with FUSE on August 29, 2003 (observation id: D0250101) and June 22, 2004 (id: D0250102) using the LWRS aperture with a resolving power of R ≈ 20 000. The individual exposure times were about 200 sec to minimize effects of orbital motion. Problems with coalignment of the telescope channels in the second observation resulted in loss of the SiC channel data and most of the LiF2 channel data; the total exposure times varied from 730 sec in the SiC channels to 2 335 sec in LiF1. The data were reduced with CalFUSE v. 3.2.2, did not result in any significant changes to the spectrum. A correction for the “worm” feature (a shadow cast by the detector grid wires) in LiF1b was obtained from a representative discussion of FUSE data reduction procedures, see Kruk et al. (2002), and Dixon et al. (2007) for a description of CalFUSE.

An additional far-UV observation was performed earlier with BEFS (R ≈ 10 000) aboard ORFEUS II on November 30, 1996 (id: BEFS2162) with an exposure time of 1 112 sec. We retrieved it from MAST. In Fig. 1 we show a comparison of ORFEUS and FUSE observations. We note that the measured flux levels agree very well. For our analysis we mainly used the FUSE observation that has better resolution and S/N. All observations are slightly smoothed with a Savitzky-Golay low-pass filter (Savitzky & Golay 1964).

The FUSE spectrum exhibits a strong contamination by interstellar line absorption (Fig. 2). To identify weak photospheric lines in the spectrum, we employed the program OWENS (cf. Lemoine et al. 2002; Hébrard et al. 2002). With OWENS, we can simulate interstellar clouds with individual parameters such as, e.g., radial velocity, column density in the line of sight, temperature of the gas, and microturbulence velocity. A large number of ions were taken into account, e.g., H i, C ii, C iii, N i, N ii, N iii, O i, Si ii, Ar i, Fe ii, and the H2 molecule (J = 0, 1, 2, 3, 4). The continuum is well matched and most of the absorption lines are well reproduced by the combined spectrum (Fig. 2). The stellar spectrum is calculated with Kurucz’ LIN lines (cf. Sect. 4) and, thus, the strong absorption feature at 979 Å (not observed) is most likely due to uncertain wavelengths in these lists.

Rauch (2000) determined an interstellar neutral hydrogen column density of nH1 = 2 · 1020 cm−2 from Ly α (IUE observation). We measured the same value from Ly β in the FUSE observation (Fig. 3). The inner line core appears much too broad at nH1 = 4 · 1020 cm−2.

In order to determine the interstellar reddening, we used FUSE and IUE observations (Fig. 4). We find that the continuum slope is best reproduced with E_B-V = 0.01 ± 0.01 us-
ing the Galactic reddening law of Seaton (1979). At this low $E_{B-V}$, we arrive at the same result if the LMC reddening law of Howarth (1983) is used. This value is significantly lower than $E_{B-V} = 0.0526$ (cf. Rauch 2000), calculated from $n_{\text{H}1}$ using an approximate formula given by Greimel & Lamers (1989).

The ISM line absorption has been neglected in this particular determination. Its consideration would decrease the flux level of our synthetic energy distribution at higher energies and, thus, the derived reddening would be even lower. Since there are uncertainties, such as in the location of the observed stellar continuum-flux level (difficult to determine in the presence of strong ISM absorption) and the validity of the applied reddening law, and the reddening is relatively small, we decided to adopt $E_{B-V} = 0.00$ for our analysis without loss of generality. Our determination of photospheric properties is not significantly affected by this assumption.

Fig. 7 shows that – even considering ISM line absorption – an additional, presently still unexplained factor, linearly increasing towards longer wavelengths, appears necessary to achieve a good fit to the continuum in the FUSE wavelength range.

3. Atomic data and modeling

The model atmosphere of the primary of LB 3459 was calculated with the Tübingen Model Atmosphere Package TMAP. Details about TMAP are given by Werner et al. (2003) and Rauch & Deetjen (2003).

In addition to H and He, the light metals C, N, O, Mg, Si, P, and S, as well as the elements of the iron group (Ca – Ni), were considered. In our final model atmosphere, 567 levels are treated in NLTE with 724 individual lines and 801 superlines.

The statistics of our final model atmospheres are summarized in Table 1. The model atoms are based on atomic data taken from the databases of NIST (National Institute of Standards and Technology), Opacity Project (Rauch et al. 1994), and Iron Project (Hummer et al. 1993). In the case of the iron-group elements, energy levels and oscillator strengths stem from Kurucz’ line lists (Kurucz 1991, see Sect. 4 for details).

Table 1. Statistics of the elements considered in our final model.

| ion | NL | L | RBB | ion | NL | RBB | RBB$_{\alpha}$ |
|-----|----|---|-----|-----|----|-----|-------------|
| H i | 15 | 1 | 105 | Ca ii | 7 | 21 | 890 |
| H ii | 1 | 0 | 0 | Ca iii | 3 | 25 | 10 241 |
| He i | 29 | 15 | 61 | Ca iv | 7 | 22 | 141 956 |
| He ii | 14 | 18 | 91 | Ca v | 7 | 24 | 65 994 |
| C ii | 1 | 41 | 0 | Sc ii | 7 | 26 | 36 087 |
| C iii | 13 | 54 | 19 | Sc iii | 7 | 24 | 67 5 |
| C iv | 14 | 44 | 35 | Sc iv | 7 | 26 | 15 024 |
| C v | 1 | 0 | 0 | Sc v | 7 | 24 | 65 994 |
| N ii | 1 | 21 | 0 | Sc vi | 1 | 0 | 0 |
| N iii | 13 | 32 | 21 | Ti ii | 7 | 26 | 154 681 |
| N iv | 34 | 60 | 108 | Ti iii | 7 | 24 | 15 722 |
| N v | 14 | 48 | 35 | Ti iv | 7 | 24 | 1 000 |
| N vi | 1 | 0 | 0 | Ti v | 7 | 24 | 26 654 |
| O ii | 2 | 3 | 0 | Ti vi | 1 | 0 | 0 |
| O iii | 19 | 35 | 33 | V iv | 7 | 24 | 37 130 |
| O iv | 11 | 58 | 19 | V v | 7 | 27 | 2 123 |
| O v | 5 | 7 | 3 | V vi | 1 | 0 | 0 |
| O vi | 1 | 0 | 0 | Cr ii | 6 | 18 | 303 129 |
| Mg i | 5 | 1 | 1 | Cr iii | 7 | 25 | 523 586 |
| Mg ii | 14 | 19 | 34 | Cr iv | 7 | 25 | 234 170 |
| Mg iii | 15 | 20 | 20 | Cr v | 7 | 24 | 43 860 |
| Mg iv | 1 | 0 | 0 | Cr vi | 1 | 0 | 0 |
| Si iii | 6 | 6 | 4 | Mn ii | 6 | 20 | 68 101 |
| Si iv | 16 | 7 | 44 | Mn iii | 7 | 22 | 671 822 |
| Si v | 1 | 0 | 0 | Mn iv | 7 | 25 | 719 387 |
| P iii | 3 | 7 | 0 | Mn v | 7 | 25 | 285 376 |
| P iv | 15 | 36 | 9 | Mn vi | 1 | 0 | 0 |
| P v | 18 | 7 | 12 | Fe ii | 6 | 20 | 218 490 |
| P vi | 1 | 0 | 0 | Fe iii | 7 | 25 | 301 918 |
| S iii | 1 | 9 | 0 | Fe iv | 7 | 25 | 1027 793 |
| S iv | 6 | 9 | 4 | Fe v | 7 | 25 | 793 718 |
| S v | 14 | 5 | 16 | Fe vi | 1 | 0 | 0 |
| S vi | 18 | 7 | 48 | Co ii | 6 | 18 | 244 873 |
| S vii | 1 | 0 | 0 | Co iii | 7 | 25 | 679 280 |
| Co iv | 7 | 25 | 552 916 |
| Co v | 7 | 25 | 1469 717 |
| Co vi | 1 | 0 | 0 |
| Ni ii | 6 | 18 | 96 647 |
| Ni iii | 7 | 22 | 418 248 |
| Ni iv | 7 | 25 | 949 506 |
| Ni v | 7 | 27 | 1 006 189 |
| Ni vi | 1 | 0 | 0 |

For iron group elements, RBB$_{\alpha}$ denotes the number of individual lines which were summed up to so-called superlines (see, e.g., Rauch & Deetjen 2003).

We adopted the parameters determined by Rauch (2000) ($T_{\text{eff}} = 42$ kK, log $g = 5.21$, and abundances) for our first models. For additional species (Ca, Sc, Ti, V, Cr, Mn, Co), we assumed solar abundances (Asplund et al. 2005). Fig. 5 shows the ionization fractions of all elements in the model atmosphere. The ionization stages in and iv dominate in the line-forming regions.

9 http://physics.nist.gov
10 http://vizier.u-strasbg.fr/topbase/topbase.html
11 http://vizier.u-strasbg.fr/tipbase/tipbase.html
Fig. 5. Temperature and density structure, as well as ionization fractions of the elements that are considered in our atmosphere model of LB 3459.

The apparent continuum slope in the FUSE spectrum changes at approximately 1100 Å, becoming noticeably shallower on the shorter-wavelength side of this point. To facilitate detailed comparison of the small-scale structure in the models with the data, the models shown in Fig. 7 have been normalized by a factor that increases linearly from $5.6 \cdot 10^{18}$ at 910 Å to $7.9 \cdot 10^{18}$ at 1190 Å.

3.1. Identification of photospheric lines

In the FUSE observation of LB 3459 (Fig. 7), several photospheric absorption lines are prominent and isolated from ISM absorption lines. In Table 2, we summarize the identified photospheric lines. For the first time, we have identified P IV and P V (Fig. 8), as well as S IV and S VI (Fig. 9), lines in the spectrum of LB 3459.}

Rauch (2000) identified a large number of strong absorption lines of Fe IV-V and Ni IV-V in IUE observations of LB 3459. In
Fig. 7. Comparison of the synthetic stellar spectrum of our final HHeCNOMgSiPS+CaScTiVCrMnFeCoNi NLTE model ($T_{\text{eff}} = 42 \, \text{kK}, \log g = 5.3$) (blue, thin line, Kurucz' POS lines) and the synthetic stellar + ISM spectrum (thick red, Kurucz LIN lines) with the FUSE observation. The observation is shifted to match the rest wavelengths of the photospheric lines marked at the top of the panels. The synthetic spectra are normalized to match the observed flux level with a factor linearly increasing from $5.6 \cdot 10^{18}$ to $7.9 \cdot 10^{18}$ within $910 \, \text{Å} – 1190 \, \text{Å}$. 
4. Spectral analysis

A detailed spectral analysis of optical and near-UV observations of LB 3459 had been performed by Rauch (2000). We considered additional species (Table 1) in order to model the far-UV spectrum reliably. Although no iron-group lines were identified in the FUSE observation, we included all iron-group elements with individual model atoms (Table 1) using the Kurucz line lists (Kurucz 1991). The model atoms and the respective atomic data files were constructed in a statistical approach (introducing "superlevels" and "superlines") with the program IroNlce (Iron Opacity Interface, Deetjen et al. 1999, Rauch & Deetjen 2003). For the model-temperature calculation, we considered all lines (so-called LIN lists that include laboratory measured, as well as computed, values) that are expected to be present in the observed spectrum.

Table 2. Identified photospheric lines in the FUSE observation of LB 3459, with only P IV 3p^2 1P^0 - 3p^2 3P lines identified, and other P IV line identifications uncertain.

| Wavelength / Å | Ion | Transition |
|----------------|-----|------------|
| 918.50         | P IV| 3p^2 1D^0 - 3d^2 1D^0 |
| 920.35         | He ii | 2 - 20 |
| 920.96         | H i  | 1 - 10  |
| 921.32         | O IV| 2p^2 1P^0 - 2p^2 3P^0 |
| 921.56         | He ii | 2 - 19 |
| 921.99 - 924.28 | N IV| 2p^2 1P^0 - 2p^2 3P^0 |
| 922.74         | He ii | 2 - 18 |
| 923.15         | H i  | 1 - 9   |
| 924.93         | He ii | 2 - 17 |
| 924.22         | S v  | 3p 1P^0 - 3p^3 1S |
| 925.81         | He i | 2 - 16  |
| 926.23         | H i  | 1 - 8   |
| 927.85         | He ii | 2 - 15 |
| 930.34         | He ii | 2 - 14 |
| 930.75         | H i  | 1 - 7   |
| 933.38 - 944.52 | Si IV| 3s^2 3S - 3p^3 1P^0 |
| 933.45         | He ii | 2 - 13 |
| 937.39         | He ii | 2 - 12 |
| 937.80         | H i  | 1 - 6   |
| 942.51         | He ii | 2 - 11 |
| 948.09 - 948.21 | C IV| 2s^2 1S - 2p^3 1P^0 |
| 948.15 - 948.61 | N IV| 3p^2 3P^0 - 3p^3 1D |
| 949.33         | He ii | 2 - 10 |
| 949.74         | H i  | 1 - 5   |
| 950.66         | P IV| 3s^2 1S^0 - 3p^3 1P^0 |
| 955.34         | N IV| 2p^2 1P^0 - 2P^2 3P |
| 958.70         | He ii | 2 - 9  |
| 967.66 - 968.18 | O iii| 3p^1 1P^0 - 3p^3 1S |
| 972.11         | He ii | 2 - 8   |
| 972.93         | H i  | 1 - 4   |
| 975.84         | S v  | 4p^2 3P^0 - 4S^2 3S |
| 977.02         | C iii | 2s^2 1S - 2p^3 1P^0 |
| 979.77 - 979.97 | N iii| 2p^2 3D - 2p^2 3D^* |
| 989.80 - 991.58 | N iii| 2p^2 3D - 2p^2 3D^* |
| 992.36         | He ii | 2 - 19 |
| 997.44 - 1000.38 | P v| 3d^2 4P^0 - 4P^2 |
| 1000.37 - 1000.54 | S v| 4d^2 3F - 5d^2 3F^* |
| 1001.75 - 1003.21 | N iii| 2p^2 3P - 3p^3 1P |
| 1005.99 - 1006.03 | N iii| 2p^2 3P - 3p^3 1P |
| 1011.00         | P IV| 3p^2 3P - 3P |
| 1025.27         | He ii | 2 - 6   |
| 1025.56         | P IV| 3p^2 3P - 3p^3 3P |
| 1025.72         | H i  | 1 - 3   |
| 1028.09 - 1035.52 | P IV| 3p^3 3P^0 - 3p^3 3P |
| 1039.92         | S v  | 3d^2 3D - 3d^2 3D^* |
| 1045.50 - 1047.27 | S v| 4p^2 3P^0 - 4d^2 3D^* |
| 1049.65         | P IV| 3d^2 3D - 3d^2 3F |
| 1050.60         | N IV| 3p^3 3P - 3p^3 1D |
| 1062.88         | S iv | 3p^3 3P - 3p^3 1D^* |
| 1064.61 - 1066.64 | P IV| 3d^2 3D - 3d^2 3D^* |
| 1073.19 - 1073.72 | S iv| 3p^3 3P - 3p^3 1D |
| 1078.71         | N iv | 3d^2 3D - 4d^2 1F |
| 1084.94         | He ii | 2 - 5   |
| 1086.09 - 1086.69 | N iv| 3p^3 3P - 3p^3 3S |
| 1086.93 - 1091.44 | P iv| 3d^2 3D - 3d^2 3F |
| 1092.35 - 1093.48 | N iv| 3d^2 3D - 3d^2 3F |
| 1097.32 - 1097.34 | C iv| 4s^2 3S - 8p^3 3P |
| 1104.53         | N iv | 3d^2 3F - 4p^3 3D |
| 1107.59 - 1107.98 | C iv| 3p^3 3F - 4d^2 3D |
| 1117.76         | S iv | 4f^2 3P - 5g^2 3G |
| 1117.93         | N iv | 3d^2 3D - 3d^2 3P |
| 1117.98 - 1128.01 | S v| 3s^2 3S - 3p^3 1P |
| 1118.55         | P iv | 3p^2 1P^0 - 3p^3 1S |
| 1122.03 - 1134.09 | S v| 3d^2 3D - 3d^2 3F |
| 1122.49 - 1128.34 | S v| 3p^2 1P^0 - 3p^2 3P |
| 1131.03 - 1132.94 | N iv| 3p^3 3P - 3p^3 3P |
| 1133.12 - 1136.27 | N iv| 3s^2 3S - 3s^2 3P |
| 1137.28         | P iv | 3d^2 3D - 3d^2 3F |
| 1149.63 - 1153.78 | O iii| 2p^2 3S - 2p^2 3P |
| 1168.48 - 1169.57 | N iv| 3d^2 3D - 3d^2 3F |
| 1168.86 - 1169.01 | C iv| 3d^2 3D - 4f^2 3F |
| 1174.93 - 1176.37 | C iii | 2p^3 1P^0 - 2p^3 3P |
| 1182.97 - 1184.57 | N iii| 2p^2 2P - 2p^2 2P |

Fig. 6. Comparison of our synthetic spectrum (thick, red line, $T_{\text{eff}} = 42$ kK, log $g = 5.3$, abundances see Fig. 14) with the IUE observation (thin line, SWP 17822, cf. Rauch 2000, his Fig. 14) of LB 3459.

Fig. 8. Comparison of synthetic spectra ($T_{\text{eff}} = 42$ kK, log $g = 5.3$, P with solar abundance) around P IV $\lambda\lambda$ 1028.09, 1030.51, 1033.11, 1035.51 Å (top) and the P V $\lambda\lambda$ 1117.98, 1128.01 Å resonance doublet (bottom) with the FUSE observation.
as theoretically calculated lines) to simulate the total opacity correctly. Figure 7 shows the good agreement of our final synthetic “LIN” spectrum with the FUSE observation. In a close inspection, e.g. around λ960 Å or λ1015 Å, opacity appears to be missing. This might be the result of partly uncertain “LIN”-line wavelengths.

For a detailed comparison with the observations, e.g. for identification and abundance determination, we have to restrict ourselves to lines measured in the laboratory (using POS lists – about 10 % of the lines in the LIN lists; Fig. 7).

4.1. Rotational velocity

The spectral analysis of Rauch (2000) was hampered by an uncertain rotational velocity. The strongest photospheric line feature in the FUSE observation of LB 3459 (Fig. 7), C III λλ 1174 – 1177 Å, is well-suited to measuring the rotational velocity (Fig. 10). We determined \( v_\text{rot} = 35 \pm 5 \text{ km/sec} \). This agrees with \( v_\text{rot} = 34 \pm 10 \text{ km/sec} \), which was calculated by Rauch (2000) in an attempt to numerically eliminate the effects of orbital smearing from their spectra with relatively long exposure times. However, Rauch (2000) and Rauch & Werner (2003) assumed this \( v_\text{rot} \) was uncertain and used \( v_\text{rot} = 45 \text{ km/sec} \) in their analyses. Our result now agrees with the \( v_\text{rot} = 34.7 - 38.7 \text{ km/sec} \) calculated under the assumption of bound rotation from the primary’s radius \( R_1 = 0.179 - 0.200 \text{ R}_\odot \) given by Hilditch et al. (2003).

From \( v_\text{rot} = 35 \pm 5 \text{ km/sec} \) and an orbital period of \( P = 22 597.033 \text{ sec} \) (Kilkenny et al. 2000), we can calculate a stellar radius of \( r_1 = 0.181 \pm 0.025 \text{ R}_\odot \). Rauch & Werner (2003) measured a radial-velocity amplitude of \( A_1 = 39.19 \pm 0.05 \text{ km/sec} \) and calculated \( a_1 = 0.2025 \pm 0.0019 \text{ R}_\odot \). With \( a_1 = 0.2025 \text{ R}_\odot \) and \( r_1 = 0.181 \text{ R}_\odot \), we can calculate \( v_\text{rot} = 31 \text{ km/sec} \), which is in good agreement. Thus, it is most likely that the rotation of LB 3459 is bound.

4.2. Effective temperature

The effective temperature \( T_{\text{eff}} = 42 \pm 1 \text{ kK} \) of LB 3459 was determined by Rauch (2000) within small error limits from the evaluation of ionization equilibria of \( \text{He I} / \text{He II}, \text{C III} / \text{C IV}, \text{N III} - \text{N V}, \text{and O IV} / \text{O V} \). In the FUSE wavelength range, we also find metal lines of successive ionization stages, namely \( \text{C III} / \text{C IV}, \text{N III} / \text{N IV}, \text{O III} / \text{O IV}, \text{P IV} / \text{P V}, \text{and S IV} / \text{S V} \). Due to the strong contamination by the ISM, many of these lines, e.g. the strong \( \text{C III} \lambda 977.03 \text{ Å} \) and \( \text{N III} \lambda 989.79 \text{ Å} \) (Fig. 7) lines, cannot be used, but there are isolated lines that are suitable for a analogous determination of \( T_{\text{eff}} \) (Figs. 8 9 11). Their ionization equilibria appear well-matched at \( T_{\text{eff}} = 42 \text{ kK} \), however, the strong contamination by ISM line absorption, the reddening, and the iron-group opacities (Sect. 5.4) make it difficult to find the continuum flux level for a proper normalization. Thus, we estimate that our \( T_{\text{eff}} \) determination cannot be better than \( \pm 3 \text{ kK} \). Consequently, we adopted \( T_{\text{eff}} = 42 \text{ kK} \) for our analysis.

4.3. Surface gravity

To investigate the impact of the rotational velocity on the determination of the surface gravity of LB 3459, we repeated the \( \chi^2 \) fit of Rauch (2000, see his Fig. 4) with the same synthetic and observed fluxes, but we used \( v_\text{rot} = 35 \pm 5 \text{ km/sec} \). We arrive now at a higher \( \log g = 5.30 (\Delta \log g = 0.09) \). These synthetic fluxes were calculated only from H+He models, so they suffer from the Balmer-line problem due to the neglecting metal opacities (Werner 1996). Consequently, we calculated a small grid of models that consider all elements from H – Ni. Due to the relatively long calculation times of these much more detailed models, we had to restrict this grid to a fixed \( T_{\text{eff}} = 42 \text{ kK}, \log g = 5.20 - 5.60 \) and \( \Delta \log g = 0.05 \). The \( \chi^2 \) fit yields \( \log g = 5.30 \) for H α and H ε and \( \log g = 5.35 \) for H β, H γ, and H δ. A comparison of the observed hydrogen Balmer series with the observations (Fig. 12) shows clearly that a higher \( \log g \) does not agree with the observations.

In the FUSE wavelength region, a higher surface gravity seems to improve the fit to the “shoulders” between the higher hydrogen Lyman lines (Fig. 13). However, the many uncertainties in the wavelengths range, e.g. due to ISM absorption, unfortunately prevent any firm conclusion.
4.4. Abundances

The possibility of determining photospheric abundances from the FUSE observation is strongly limited by both the ISM line absorption (Fig. 7) and the limited number of Kurucz' POS lines (Sect. 4). Only a few isolated photospheric lines are accessible in this wavelength range. We adopted the previously determined abundances given by Rauch (2000) and the adjusted Si and Ni abundances (Sect. 3), because test calculations have shown that these give a good fit to the observation. The newly introduced elements phosphorus and sulfur enabled us to identify lines of P IV and P V (Fig. 8) and S IV and S VI (Fig. 9). The determined abundances are solar and 0.01 times solar, respectively. Only weak lines of the iron-group element are located in the FUSE wavelength range. Due to the stellar rotation (Sect. 4.1), we were not able to identify any of these lines.

5. Results and conclusions

We performed an NLTE spectral analysis of FUSE observations of the post common-envelope binary LB 3459. The short FUSE exposure times (200 sec) allowed us to measure the rotational velocity $v_{\text{rot}} = 35 \pm 5$ km/sec of the primary star of LB 3459. This is in good agreement with the $v_{\text{rot}} = 34.7 - 38.6$ km/sec given by Hilditch et al. (2003), and it confirms that the rotation of LB 3459 is bound.

A re-analysis of optical spectra (cf. Rauch 2000) has shown that log $g$ is about 0.1 higher than given by Rauch (2000, log $g=5.2$), who assumed bound rotation because the system has been classified to be a post common-envelope binary consisting of an sdOB and a main-sequence star (de Kool & Ritter 1993) where the common-envelope phase is much longer than the synchronization time. Since the primary’s radius was as-
The “gravity problem” is still unsolved although the surface gravity is slightly higher: \( \log g = 5.3 \) rather than the 5.2 found by Rauch (2000). If we take the radius of the primary as \( r_1 = 0.236 R_\odot \) due to the lower \( \log g \), a higher \( v_{\text{rot}} = 45 \text{ km/sec} \) was adopted then for the analysis. From \( v_{\text{tot}} = 35 \pm 5 \text{ km/sec} \) and an orbital period of \( P = 22,597.033 \text{ sec} \) (Kilkenny et al. 2000), we can calculate a stellar radius of \( r_1 = 0.181 \pm 0.025 R_\odot \).

The spectral analysis is hampered by a strong ISM contamination. We used OWENS to model the ISM line absorption qualitatively in order to identify and model stellar lines. Our NLTE models consider opacities of 18 species from H – Ni. It is obvious that the iron-group elements (here: Ca – Ni) contribute to the opacity. However, we were not able to identify any individual iron-group line in the FUSE observation.

We identified phosphorus and sulfur lines in the FUSE spectrum. For phosphorus we can determine a solar abundance. The sulfur abundance is surprisingly low, and we determined a 0.01 times solar abundance. The photospheric abundances of LB 3459 are summarized in Fig. 14. The derived abundance pattern reflects the interplay of gravitational settling and radiative levitation (cf. Rauch 2000).

\[
\begin{align*}
\log (\text{mass fraction}\, / \, \text{solar mass fraction}) &\approx g \approx 5.5 \\
\text{where G is the gravitational constant. Note that log g \approx 5.5 is necessary for good agreement (cf. [Rauch 2000])}
\end{align*}
\]

An additional uncertainty (cf. Rauch 2000) is the lack of appropriate evolutionary models for post-common-envelope binaries to compare so as to derive the primary’s mass.

High-resolution and high S/N observations in the near-UV wavelength range is highly desirable when searching for the weak lines of iron-group elements. Figure 15 demonstrates that lines of all elements from Ca – Ni but Sc and V should be detectable. The search for signatures of the “heated-up” secondary in IR observations in order to further investigate the nature of this star also appears possible, and its contribution to the continuum at \( \lambda = 12 \mu \text{ is a few percent.} \)

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