THE STRUCTURE OF THE LOCAL INTERSTELLAR MEDIUM. V. ELECTRON DENSITIES

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ABSTRACT

We present a comprehensive survey of C ii absorption detections toward stars within 100 pc in order to measure the distribution of electron densities present in the local interstellar medium (LISM). Using high spectral resolution observations obtained by GHRS and STIS on board HST, we searched for all detections of LISM C ii absorption. We identify 13 sight lines with 23 individual C ii absorption components, which provide electron density measurements. We employ several strategies to determine more accurate C ii column densities from the saturated C ii resonance line, including constraints of the line width from the optically thick C ii line, constraints from independent temperature measurements of the LISM gas based on line widths of other ions, and third, using measured S ii column densities as a proxy for C ii column densities. The distribution of electron densities based on using S ii as a proxy for C ii is similar to the distribution based on carbon alone, while significantly tighter, and proves to be a promising technique to avoid grossly overestimating the C ii column density based on the saturated line profile. The sample of electron densities appears consistent with a lognormal distribution and an unweighted mean value of \( n_e (\text{C } ii) = 0.11_{-0.05}^{+0.10} \text{ cm}^{-3} \). Seven individual sight lines probe the Local Interstellar Cloud (LIC), and all present a similar value for the electron density, with a weighted mean of \( n_e (\text{LIC}) = 0.12 \pm 0.04 \text{ cm}^{-3} \) Given some simple assumptions, the range of observed electron densities translates into a range of thermal pressures, \( P/k = 3300^{+5500}_{-1900} \text{ K cm}^{-3} \). This work greatly expands the number of electron density measurements and provides important constraints on the ionization, abundance, and evolutionary models of the LISM.

Subject headings: atomic processes — ISM: abundances — line: profiles — solar neighborhood — techniques: spectroscopic — ultraviolet: ISM

1. INTRODUCTION

More than a decade of use of high spectral resolution (\( R \equiv \lambda/\Delta \lambda \geq 50,000 \)) ultraviolet (UV) spectrographs with wide spectral coverage, such as the Goddard High Resolution Spectrograph (GHRS) and particularly the Space Telescope Imaging Spectrograph (STIS) on board the Hubble Space Telescope (HST), has been a boon for observations of the local interstellar medium (LISM). Since the LISM, the collection of warm gas in the immediate (<100 pc) vicinity of the Sun, is traversed by all sightlines that extend beyond our solar system, the number of observations with LISM absorption has grown tremendously. The use of this expanded database has been the impetus for many investigations into the nature of the LISM, including the other papers in this series, which present fits to LISM absorption in many ions (e.g., D i, C ii, N i, O i, Mg ii, Al ii, Si ii, and Fe ii; Redfield & Linsky 2002, 2004a), temperature and turbulent velocity measurements in LISM clouds (Redfield & Linsky 2004b), and a dynamical model of the LISM (Redfield & Linsky 2008).

Due to the low column densities [log \( N(\text{H i}) \sim 16.8-18.3 \)] typical of local clouds, the ionization structure of hydrogen is vital to understanding the physical structure and origins of the LISM. An accurate accounting of the ionized fraction of hydrogen (and the ionization levels of all ions) is critical for measuring abundances and the depletion of gas-phase ions onto dust grains. Low column density clouds [i.e., log \( N(\text{H i}) < 19.5 \)], like the LISM clouds, are not significantly shielded to ionizing photons (Jenkins 2004; Jenkins et al. 2000b; Sofia & Jenkins 1998). The origin and evolution of the Local Bubble, the ~100 pc radius cavity in which the warm LISM clouds reside (Lallement et al. 2003), are encoded in the ionization structure. Recent nondetections of high-temperature lines in the extreme-UV (EUV; Hurwitz et al. 2005) and the realization that soft X-ray emission caused by the heliosphere (Lallement 2004) may contribute to the emission formerly assigned to nearby hot gas (Snowden et al. 1990) highlight the current challenges in understanding the thermal structure and ionization level of the Local Bubble. A realistic Local Bubble can be modeled far from ionization equilibrium (Breitschwerdt & de Avillez 2006), and the physical structure of the Local Bubble can have a direct influence on the ionization structure of the warm LISM clouds (Slavin & Frisch 2002). In addition, the structure of the heliosphere, the interface between the LISM and the solar wind, depends significantly on the ionization structure of the surrounding ISM (Müller et al. 2006).

The ionization fraction of hydrogen, \( X(\text{H}) = n(\text{H})/(n(\text{H}) + n(\text{H} \text{ii})) \), is computed using \( n(\text{H}) = n(\text{He} \text{ii}) \times N(\text{H} \text{ii})/N(\text{He} \text{ii}) \) and \( n(\text{H} \text{ii}) \approx n_e \). Typically, \( n(\text{He} \text{ii}) \) is derived from in situ measurements of interstellar helium streaming through the heliospheric interface into the inner solar system (Gloeckler et al. 2004), which, of course, is only a measurement of the helium density at one point in the LISM. \( N(\text{H} \text{ii})/N(\text{He} \text{ii}) \) is derived from EUV observations of nearby white dwarfs (WDs), which contain the ionization edge of He i, and the continuum provides an estimate of H i (Dupuis et al. 1995; Barstow et al. 1997). The weighted mean from this work, based on nine WD sight lines within 100 pc, is \( N(\text{H} \text{ii})/N(\text{He} \text{ii}) = 14.1 \pm 1.7 \), although this is a full sight-line average, since this technique cannot separate individual cloud components.

The remaining measurement in the calculation of the ionization structure is the electron density, \( n_e \). The methods employed to make this measurement utilize atomic transitions along interstellar...
sight lines, which have the benefit of being able to resolve individual absorbers if taken at high spectral resolution and can provide a large number of measurements through various LISM environments. The ratio of magnesium ionization stage column densities, $N$(Mg ii)/$N$(Mg i), has provided a number of $n_e$ measurements (e.g., Frisch et al. 1990; Lallement et al. 1994; Frisch 1994; Lallement & Ferlet 1997). However, this technique suffers from the requirement of ionization equilibrium and a strong temperature dependence. Alternatively, the ratio of the collisionally excited carbon line column density to the resonance line column density, $N$(C ii)/$N$(C i), can provide $n_e$ estimates without the need for ionization equilibrium and has a very weak temperature dependence. However, due to the weakness of the excited C ii absorption and saturation of the only available UV C ii resonance line, few LISM sight lines have been analyzed using this technique. Wood & Linsky (1997) used high spectral resolution ($R \sim 100,000$) spectra of $\alpha$ Aur to calculate $n_e$ along the line of sight and demonstrated that the relatively simple absorption profiles through the LISM provide an excellent opportunity to make precise measurements of the electron density. Holberg et al. (1999) used the same technique to measure the electron density in ISM absorbers along the line of sight toward WD 1029+537, a WD $\sim$132 pc away. The excited absorption along other LISM sight lines has been measured at moderate resolution ($R \sim 20,000$) with Copernicus (e.g., York & Kinahan 1979) and the Far Ultraviolet Spectroscopic Explorer (FUSE; Lehner et al. 2003), although typically individual absorbers are not resolved, so the estimates are full sight-line averages. A few $n_e$ measurements in slightly warmer, and possibly local, O vi−bearing gas, derived by comparisons of O vi emission (Dixon et al. 2006) and absorption (Lehner et al. 2003), have led to $n_e$ estimates of $\sim$0.2 cm$^{-3}$.

We present an inventory of high spectral resolution C ii detections along LISM sight lines from the complete HST spectroscopic database. We employ various strategies to circumvent one of the challenges of using the C ii/C iii ratio technique, namely, obtaining a reliable C ii column density from a saturated resonance line: (1) simultaneously fitting the optically thin C ii and saturated C iii profiles puts realistic constraints on the Doppler width and column densities of the resonance line; (2) use of previously determined LISM temperatures and turbulent velocities, typically derived from the same data set (Redfield & Linsky 2004b), constrains the Doppler width of the carbon lines; and (3) use of optically thin profiles of S ii, which has a similar ionization potential as C ii, as a proxy of the C ii column density (e.g., Oliveira et al. 2003). This collection of new, high spectral resolution UV observations of nearby stars presents an opportunity to greatly expand the number of LISM electron density measurements and probe the ionization structure of our most immediate interstellar environment.

2. OBSERVATIONS

For our purpose, we are interested in sight lines toward stars within 100 pc that show interstellar absorption in both C ii ($\lambda$1335.5323) and the C ii doublet ($\lambda$1335.6627 and 1335.7077).

We compiled all moderate- to high-resolution observations of nearby stars with the HST spectrographs: GHRS and STIS. The complete sample includes 417 unique targets within 100 pc, almost half of which have spectra that cover the wavelength region of C ii and C iii. All these relevant spectra were scrutinized for signs of C ii absorption. We found only 13 sight lines that show LISM absorption in both transitions and list them in Table 1.

The 100 pc distance limit is chosen to coincide with the approximate extent of the Local Bubble (Lallement et al. 2003). Observations of more distant stars may be more difficult to analyze, as they are more likely to traverse many absorbing clouds creating a blended line profile.

Table 2 lists the observational parameters of the data sets extracted from the HST Data Archive. All observations are necessarily moderate ($R \gtrsim 20,000$) to high ($R \gtrsim 100,000$) spectral resolution in order to resolve narrow, closely spaced interstellar absorption features and to increase the likelihood of the detection of the weak C ii line.

We analyzed LISM absorption in the S ii multiplet ($\lambda$1250.578, 1253.805, and 1259.518) in order to better constrain the C ii column density. Ten of the 13 targets have spectra that include the S ii lines. In addition, we fit the LISM absorption in Mg i ($\lambda$2852.9631) and the Mg ii doublet ($\lambda$2796.3543 and 2803.5315) in order to further constrain the electron density. Only 3 of the 13 targets have spectra that include both Mg i and Mg ii. Observations utilized for these analyses are also included in Table 2.

We reduced the GHRS data acquired from the HST Data Archive with the CALHRS software package using the Image Reduction

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**TABLE 1**

| HD       | Star Name | Other Name | Spectral Type | $l$ (deg) | $b$ (deg) | Distance (pc) |
|----------|-----------|------------|---------------|-----------|-----------|---------------|
| 34029    | $\alpha$ Aur | Capella | G0 III+G8 III | 162.58 | +4.566 | 12.94 ± 0.15 |
| 120315   | $\eta$ UMa | Alcald | B3 V | 100.69 | +65.32 | 30.87 ± 0.71 |
| 209952   | $\alpha$ Gru | Alnair | B7 IV | 349.99 | −52.47 | 31.10 ± 0.79 |
| EX Hya   |           |           | M5/M6+DA | 303.18 | +33.62 | 64.5 ± 1.2 |
| 6457     | 74 Psc B |           | A0 V | 127.34 | −41.28 | 68.2 ± 0.8 |
| 6456     | 74 Psc A |           | A1 V | 155.95 | +7.099 | 69 ± 15 |
| WD 0501+527 |           | G191-B2B | DA1b | 165.96 | −50.26 | 74 ± 20 |
| WD 0232+035 |           | Feige 24 | DAZQO1 | 87.262 | −45.11 | 79 |
| WD 2309+105 |           | GD 246 | DA1b | 135.61 | +63.11 | 87 |
| 128345   | $\rho$ Lup |           | B5 V | 302.13 | +9.857 | 95.1 ± 6.4 |
| IX Vel   |           |           | B8 V+DA | 264.92 | −7.890 | 96.3 ± 9.1 |

**Note.**—All values are from SIMBAD unless otherwise noted.

* All distances are Hipparcos distances (Perryman et al. 1997) except for WD 0050−332 and WD 2309+105, whose distances are from Vennes et al. (1997), and WD 1210+533, whose distance is from Holberg et al. (1998).

b WD spectral types taken from McCook & Sion (1999).
and Analysis Facility (IRAF; Tody 1993) and the Space Telescope Science Data Analysis System (STSDAS). We used the most recent reference calibration files. Many of the echelle observations were obtained in the FP-SPLIT mode to reduce fixed pattern noise. The individual readouts of the FP-SPLIT spectra are combined using a cross-correlation procedure called HRS_MERGE (Robinson et al. 1992). The reduction included assignment of wavelengths using calibration spectra obtained during the course of the observations. The calibration spectra include either a WAVECAL, a direct Pt-Ne lamp spectrum used to derive the dispersion relation, or a SPYBAL (Spectrum Y-Balance), which only provides a zero-point wavelength offset. Any significant errors involved in the wavelength calibration are included in our central velocity determinations. The wavelength calibration of the Fe ii spectrum of α Gru is used to calibrate the wavelength solution of C ii, since no wavelength calibration of this segment of the spectrum was taken at the time.

We reduced the STIS data acquired from the HST Data Archive using the STIS team’s CALSTIS software package written in IDL (Lindler 1999). The reduction included assignment of wavelengths using calibration spectra obtained during the course of the observations. We used the ECHELLE_SCAT routine in the CALSTIS software package to remove scattered light. However, the scattered light contribution is negligible in this spectral range and does not influence the uncertainties in our spectral analysis.

### TABLE 2

| Star Name          | Other Name | Instrument | Grating | Spectral Range (Å) | Resolution (λ/Δλ) | Exposure Time(s) | Data Set |
|--------------------|------------|------------|---------|--------------------|-------------------|------------------|----------|
| α Aur.............................. | Capella    | GHR   | ECH-A   | 1331–1338          | 100000            | 761.6            | ZUW030BT  |
| η UMa.............................. | Alcald     | GHR   | ECH-A   | 1332–1339          | 100000            | 1171.4           | Z3CL020ET  |
|                    | Alcald     | GHR   | ECH-B   | 2790–2805          | 100000            | 108.8            | Z3CL030HT  |
|                    | Alcald     | GHR   | ECH-B   | 2845–2859          | 100000            | 652.8            | Z3CL0308T  |
|                    | Alcald     | ...   | ...     | ...                | ...              | ...              | Z3CL0307T  |
| α Gru.............................. | Alnair     | GHR   | G160M   | 1309–1345          | 20000             | 230.4            | Z1720109T  |
|                    | Alnair     | GHR   | ECH-B   | 2791–2806          | 100000            | 108.8            | Z1720209T  |
|                    | Alnair     | GHR   | ECH-B   | 2847–2861          | 100000            | 217.6            | Z172020AT  |
| WD 0050–332 .......... | GD 659     | STIS  | E140H   | 1170–1372          | 114000            | 4134             | O4GI01010 |
|                    | GD 659     | STIS  | E140M   | 1140–1735          | 45800             | 15200            | O68301010 |
|                    | EX Hya     | ...   | E140M   | ...                | ...              | ...              | O68301020 |
|                    | ...        | ...   | ...     | ...                | ...              | ...              | O68301030 |
|                    | ...        | ...   | ...     | ...                | ...              | ...              | O68302010 |
|                    | ...        | ...   | ...     | ...                | ...              | ...              | O68302020 |
|                    | ...        | ...   | ...     | ...                | ...              | ...              | O68302030 |
| 74 Psc B...................... | HR 311     | STIS  | E140H   | 1170–1372          | 114000            | 2768             | O56L02010 |
|                    | HR 311     | ...   | ...     | ...                | ...              | ...              | O56L02020 |
|                    | WD 0501+527 | ...   | ...     | ...                | ...              | ...              | O56L02030 |
|                    | G191-B2B   | STIS  | E140H   | 1170–1517          | 114000            | 5623             | O57U01040 |
|                    | G191-B2B   | ...   | ...     | ...                | ...              | ...              | O6HB01040 |
|                    | G191-B2B   | ...   | ...     | ...                | ...              | ...              | O6HB01050 |
|                    | G191-B2B   | ...   | ...     | ...                | ...              | ...              | O6HB01060 |
|                    | G191-B2B   | ...   | ...     | ...                | ...              | ...              | O6HB01070 |
|                    | G191-B2B   | ...   | ...     | ...                | ...              | ...              | O6HB01080 |
|                    | G191-B2B   | ...   | ...     | ...                | ...              | ...              | O6HB01090 |
|                    | G191-B2B   | STIS  | E230H   | 2624–3095          | 114000            | 6013             | O6HB30080 |
|                    | G191-B2B   | ...   | ...     | ...                | ...              | ...              | O6HB30090 |
|                    | G191-B2B   | ...   | ...     | ...                | ...              | ...              | O6HB300B0 |
|                    | G191-B2B   | ...   | ...     | ...                | ...              | ...              | O6HB300C0 |
|                    | G191-B2B   | ...   | ...     | ...                | ...              | ...              | O6HB300D0 |
|                    | G191-B2B   | ...   | ...     | ...                | ...              | ...              | O6HB300E0 |
| 74 Psc A...................... | HR 310     | STIS  | E140H   | 1170–1372          | 114000            | 4128             | O56L01010 |
|                    | HR 310     | ...   | ...     | ...                | ...              | ...              | O56L01020 |
|                    | HR 310     | ...   | ...     | ...                | ...              | ...              | O56L01030 |
|                    | WD 0232+035 | ...   | ...     | ...                | ...              | ...              | O56L10100 |
|                    | Feige 24   | STIS  | E140M   | 1150–1735          | 45800             | 4176             | O4G701010 |
|                    | Feige 24   | ...   | ...     | ...                | ...              | ...              | O4G702010 |
|                    | WD 2309+105 | ...   | ...     | ...                | ...              | ...              | O4G102020 |
|                    | WD 1210+533 | ...   | ...     | ...                | ...              | ...              | O5F203010 |
|                    | ...        | ...   | ...     | ...                | ...              | ...              | O5F203020 |
|                    | ...        | ...   | ...     | ...                | ...              | ...              | O5F204010 |
|                    | ...        | ...   | ...     | ...                | ...              | ...              | O5B02010 |
|                    | ...        | ...   | ...     | ...                | ...              | ...              | O5B03010 |

a For targets with multiple data sets, the exposure time listed is the sum of the exposure times for each data set.
b Spectral range encompassed by multiple data sets.
3. DATA ANALYSIS AND LINE PROFILE FITTING

Figures 1a and 1b show the C \textsc{ii} resonance and the C \textsc{ii}' excited absorption lines observed toward the 13 targets listed in Table 1. The spectra (histograms) are plotted in heliocentric velocity. Also plotted are the individual interstellar component fits (dashed lines) and the total interstellar absorption convolved with the instrumental profile (thick solid lines). As explained in § 3.2, the resonance-line plots come from the simultaneous fits, and the excited-line plots come from the individual excited-line fits.

The stellar continuum level (thin solid lines) illustrates our flux estimates of the spectra without any interstellar absorption. These continua were determined by fitting polynomials to the spectra have relatively smooth continua, and we are able to identify features, placing the continuum can be a difficult task. All of our target spectra have relatively smooth continua, and we are able to easily reproduce the intrinsic continuum flux backgrounds with the use of information retained in the, albeit saturated, resonance line. It is made evident on the WD spectra, although their broad emission lines allowed for simple continuum placement.

We used a Gaussian ISM absorption profile fitting algorithm to determine the fits. The program makes use of rest wavelengths and oscillator strengths from Morton (2003), and the instrumental line-spread functions for GHRS and STIS spectra are taken from Redfield & Falcon 3.1. Fit Parameters

Table 3 lists the fit parameters measured, including central velocity (\(v_0\) [km s\(^{-1}\)], Doppler width (\(b\) [km s\(^{-1}\)]), and the logarithm of the column density (\(N_\text{cm}^{-2}\)). The central velocity is the mean radial velocity of the absorbing component. The Doppler width is a function of the temperature (\(T\) [K]) and turbulent velocity (\(\xi\) [km s\(^{-1}\)]) of the interstellar gas:

\[
b^2 = \frac{2kT}{m} + \xi^2 = 0.016629 \frac{T}{A} + \xi^2, \tag{1}
\]

where \(k\) is Boltzmann's constant, \(m\) is the mass of the observed ion, and \(A\) is the atomic mass (\(A_{\text{C}} = 12.011\)). The column density is a measure of the amount of material along the line of sight to the target.

3.2. Procedure

For each target, we first fit the excited line alone. Since the excited line is actually a doublet, each paired absorption line possesses the same intrinsic central velocity, Doppler width, and column density, so these parameters are constrained accordingly to be identical for both lines of the doublet. We then fit the excited (C \textsc{ii}') and resonance (C \textsc{ii}) lines simultaneously, requiring the central velocities and Doppler widths to be identical. Since the resonance line is typically saturated, \(N(C \textsc{ii})\) is not well constrained, but fitting the lines simultaneously allows us to make use of information retained in the, albeit saturated, resonance line, while retrieving vital constraints from the optically thin excited line. Thus, the final fit parameters for the resonance line are derived from the simultaneous fits only. For the excited-line fits, we supplement the results from the simultaneous fits with those from the individual fits in order to limit the impact of systematic errors; these values are weighted means of the parameters from both the simultaneous fits and individual excited-line fits.

In saturated lines, Doppler width and column density are tightly coupled. Thus, any available constraints on the \(b\)-value will yield a more certain measurement for \(N\), which in the case of C \textsc{ii} and C \textsc{ii}' is critical in determining electron densities. Consequently, we use measured LISM temperatures and turbulent velocities from Redfield & Linsky (2004b) to constrain the Doppler widths for the \(\alpha\) Aur, \(\eta\) UMa, and G191-B2B lines of sight. Unfortunately, LISM temperatures and turbulent velocities have been determined for only a limited number of sight lines, so we are able to include this type of constraint for only these three targets.

Figure 1 shows four WDs (WD 0050−332, WD 0232+035, WD 2309+105, and WD 1210+533) that each possess a resonance-line component that does not have an excited-line counterpart. In each case, the Doppler width and column density for the resonance component are largely determined by what is not accounted for by the other paired components; the C \textsc{ii} measurements of the component without a C \textsc{ii}' counterpart are poorly determined as they lack the critical constraints provided by the optically thin excited component. Upper limits (3 \(\sigma\)) are determined for the C \textsc{ii}' column density for these components that show absorption only in the resonance line.

The dominant source of systematic error, as mentioned earlier, is the saturation of the C \textsc{ii} resonance line. It is made evident on comparing our results for WD 0232+035 with those from Vennes et al. (2000). Using the same data sets, Vennes et al. (2000) observed two interstellar components toward WD 0232+035 with central velocities of 3.1 \(\pm\) 0.2 and 17.6 \(\pm\) 0.9 km s\(^{-1}\), in good agreement with our measurements of 3.81 \(\pm\) 0.47 and 17.4 \(\pm\) 4.6 km s\(^{-1}\). Their measured column density of log \(N(C \textsc{ii}) = 14.32 \pm 0.57\) is a factor of 7 lower than our measurement of log \(N(C \textsc{ii}) = 15.18 \pm 0.42\), but both have large error bars, and the difference is just over 1 \(\sigma\). In contrast, the resonance-line component that has no accompanying excited-line component is measured by Vennes et al. (2000) with a column density log \(N(C \textsc{ii}) = 14.16 \pm 0.41\), a factor of 60 lower than our measured value of log \(N(C \textsc{ii}) = 15.96^{+0.32}_{-0.55}\), and a difference of almost 3 \(\sigma\). The critical factor in the disagreement concerns the treatment of the Doppler width (\(b\)). Vennes et al. (2000) allow for a range of \(b\)-values (3.5−10.0 and 4.0−10.0 km s\(^{-1}\) for the \(\sim3\) and \(\sim17\) km s\(^{-1}\) components, respectively), while we allow \(b\) to remain an independent variable. Such a wide range of \(b\)-values appears to be unnecessary; for our C \textsc{ii} fits to all sight lines, the weighted mean Doppler width is \(b = 4.30 \pm 0.87\) km s\(^{-1}\) and inconsistent with a large \(b\)-value for WD 0232+035. In addition, our fit to the S \textsc{ii} lines, which also shows two components at similar velocities (4.11 \(\pm\) 0.64 and 16.8 \(\pm\) 3.8 km s\(^{-1}\)), has a weighted mean Doppler width of \(b = 3.95 \pm 0.79\) km s\(^{-1}\), also inconsistent with a large Doppler width.

Simultaneously fitting the unsaturated excited line with the accompanying saturated resonance line is the critical property of our method, which allows us to more tightly constrain the observed column density of the saturated line than if we were fitting it alone. Essentially, the optically thin excited line provides the constraint on the Doppler width, which thereby constrains the acceptable range of column densities.

3.3. Component Determination

In determining the number of absorption components, we use the fewest number of components that produce a satisfactory fit. In most cases, the resolution is high enough and the differences between component velocities are great enough such that discernment of components is almost trivial. Other times, however, it is not so easily determined. Components are added if a clear asymmetry is detected in either the C \textsc{ii} or C \textsc{ii}' profiles (e.g., \(\eta\)
Fig. 1.—Fits of the interstellar C ii (λ1334.5) and C iii (λλ1335.6627 and 1335.7077) absorption toward 13 nearby stars. The ratios of the column densities are used to estimate the electron density for each component. The name of the target star is given above each group of plots, and the wavelength (in Å) of each line is provided within each plot. Both the 1335.6627 and 1335.7077 Å C iii lines are shown in the bottom plot. Although the 1335.6627 Å line is weak, it is evident as the blueward component in some spectra (e.g., G191-B2B, 74 Psc A and B, and ρ Lup). The data are shown in histogram form. The thin solid lines are our estimates of the intrinsic stellar flux across the absorption feature. The dashed lines are the best-fit individual absorption lines before convolution with the instrumental profile. The thick solid line represents the combined absorption fit after convolution with the instrumental profile. The spectra are plotted vs. heliocentric velocity. The parameters for these fits are listed in Table 3.
UMa, WD 1210+533, IX Vel), or if other ions indicate that additional components are required (e.g., α Gru, ρ Lup). Our general attitude toward this part of the analysis is to approach the absorption feature with Ockham’s razor in hand, so as to cut away any unjustified components.

### 3.4. S ii

We are motivated to find an alternative means of estimating the C ii column density, due to the difficulty of obtaining it directly from the strongly saturated resonance line. We use the optically thin S ii triplet, located near the C ii lines, as a proxy for C ii due to their similarity in ionization potential, which is 24.4 eV for C ii and 23.3 eV for S ii, and their similarity in the ratio of their ionization and recombination rates (defined as $P$ in Sofia & Jenkins [1998], where S ii and C ii, among other ions, are compared in their Fig. 3). Ten of the 13 sight lines with C ii absorption also have spectra that cover the S ii wavelength range. We fit all three S ii lines simultaneously, and due to the range of opacities of the three optically thin lines, we are able to measure an accurate S ii column density. The spectra and fits are shown in Figures 2a and 2b, and the fit parameters are listed in Table 4.

In order to convert from $N$(S ii) to $N$(C ii), we need to take into account the different abundances and depletion levels of sulfur and carbon in the LISM:

$$N(C_{ii_{S}}) = N(S_{ii}) \times 10^{[(C_{o} + D(C)) - [S_{o} + D(S)]]},$$

(2)

where $N(C_{ii_{S}})$ is the estimated C ii column density based on S ii as a proxy, $N(S_{ii})$ is our measured column density of S ii, $C_{o} = 8.39 \pm 0.05$ and $S_{o} = 7.14 \pm 0.05$ are the solar abundances of carbon and sulfur (Asplund et al. 2005), and $D(C)$ and $D(S)$ are the depletion levels (Jenkins 2004). Here $D(S) \sim 0.0$, as has been detected in local and more distant ISM (Lehner et al. 2003; Welty et al. 1999), and $D(C) \sim -0.17 \pm 0.19$ from Jenkins (2004), where the error incorporates some of the natural variation in the ISM [e.g., $D(C) \sim 0.0$, Cardelli et al. 1996; $D(C) \sim -0.2$, Lehner et al. 2003; $D(C) \sim -0.4$, Welty et al. 1999].
### Table 3

**Fit Parameters for C II ISM Velocity Components**

| Star Name | Other Name | Component Number | \( v \) (km s\(^{-1}\)) | \( b \) (km s\(^{-1}\)) | \( \log [N(C II)/\text{cm}^2] \)^a | \( \log [N(C II)/\text{cm}^2] \)^b | \( \log [N(C II)/\text{cm}^2] \) |
|-----------|------------|------------------|------------------------|------------------------|--------------------------|--------------------------|------------------------|
| \( \alpha \) Aur | Capella | 1 | 20.78 ± 0.28 | 3.48 ± 0.15 | 14.67 ± 0.14 | ... | 12.62 ± 0.07 |
| \( \eta \) UMa | Alcincid | 1 | -2.09 ± 0.24 | 3.76 ± 0.41 | 14.43 ± 0.12 | ... | 13.83 ± 0.12 |
| \( \alpha \) Gru | Alnair | 2 | 9.72 ± 0.95 | 5.60 ± 0.29 | 13.12 ± 0.06 | ... | 11.60 ± 0.05 |
| \( \alpha \) Gru | Alnair | 2 | -22.50 ± 0.26 \( ^a \) | 3.35 ± 0.73 | 14.98 ± 0.19 | ... | 12.69 ± 0.14 |
| \( \alpha \) Gru | Alnair | 2 | -12.63 ± 0.26 \( ^b \) | 3.49 ± 0.51 | 14.52 ± 0.23 | ... | 12.73 ± 0.18 |
| \( \alpha \) Gru | Alnair | 3 | -8.32 ± 0.60 \( ^b \) | 3.9 ± 2.2 | 13.63 ± 0.23 | ... | 12.17 ± 0.28 |
| WD 0050–332 | GD 659 | 1 | 6.17 ± 0.38 | 4.31 ± 0.77 | 15.12 ± 0.32 | ... | 14.808 ± 0.05 \( ^{0.216} \) | 12.77 ± 0.05 |
| WD 0050–332 | GD 659 | 2 | 12.3 ± 1.0 | 4.67 ± 0.67 | 15.12 ± 0.29 | ... | 14.062 ± 0.30 \( ^{0.046} \) | <11.9 |
| EX Hya | ... | ... | ... | ... | ... | ... | ... |
| 74 Psc B | HR 311 | 2 | -5.75 ± 0.50 | 4.00 ± 0.93 | 16.03 ± 0.16 | ... | 15.444 ± 0.07 \( ^{0.127} \) | 13.07 ± 0.02 |
| WD 0501+527 | G191-B2B | 1 | 5.98 ± 0.17 | 4.10 ± 0.29 \( ^b \) | 14.81 ± 0.31 | ... | 14.345 ± 0.07 \( ^{0.224} \) | 13.14 ± 0.02 |
| 74 Psc A | HR 310 | 2 | -5.97 ± 0.52 | 3.82 ± 0.59 | 16.42 ± 0.14 | ... | 15.399 ± 0.06 \( ^{0.210} \) | 13.07 ± 0.03 |
| WD 0232+035 | Feige 24 | 1 | 3.81 ± 0.47 | 4.36 ± 0.53 | 15.18 ± 0.42 | ... | 14.877 ± 0.07 \( ^{0.224} \) | 13.00 ± 0.02 |
| WD 2309+105 | GD 246 | 2 | 17.4 ± 4.6 | 3.1 ± 1.4 | 15.96 ± 0.55 | ... | 13.812 ± 0.30 \( ^{0.412} \) | <11.8 |
| WD 1210+533 | GD 246 | 2 | -9.7 ± 2.5 \( ^a \) | 4.20 ± 0.34 | 15.16 ± 0.30 | ... | 15.321 ± 0.05 \( ^{0.209} \) | 13.05 ± 0.04 |
| \( \rho \) Lup | HR 5453 | 1 | -16.11 \( ^f \) | 5.17 ± 0.74 | 15.80 ± 0.45 | ... | 14.632 ± 0.04 \( ^{0.336} \) | 13.00 ± 0.16 |
| \( \chi \) Vel | HR 5453 | 2 | -9.1 \( ^f \) | 4.3 ± 1.1 | 15.14 ± 0.61 | ... | 15.012 ± 0.30 \( ^{0.172} \) | 13.08 ± 0.11 |
| 3.5. Mg I and Mg II | ... | ... | ... | ... | ... | ... | ... |

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\( ^a \) Excited-line parameters derived from simultaneous fits.

\( ^b \) Excited-line parameters are weighted means of simultaneous and excited-only fits.

\( ^c \) Fixed \( b \)-values based on independent temperature and turbulent velocity measurements (Redfield & Linsky 2004b).

\( ^d \) Fixed velocity difference between components based on measurements of interstellar Fe II and Mg II (Redfield & Linsky 2002).

\( ^e \) Fixed velocity difference between components based on measurements of interstellar Si II.

\( ^f \) Fixed velocities based on measurements of interstellar S II.

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The C II column density estimated using this technique is listed in Table 3, and the comparison with the direct C II column density based on the C II and C II' fits shown in Figures 1a and 1b is exhibited in Figure 3. We present two errors based on this calculation: the first simply propagates the error in \( N(S II) \), while the second includes the errors in the solar abundances (Asplund et al. 2005) and the natural variation of depletions (Jenkins 2004). Figure 3 demonstrates that S II is a reasonable proxy for C II, where 12/20 (60%) agree within 2 \( \sigma \), and those that disagree by >2 \( \sigma \) are all overestimated by measuring the strongly saturated C II resonance line directly. Forcing \( N(C II) = N(C II') \) in fitting C II and C II' produces viable fits to the data that are indistinguishable from those shown in Figures 1a and 1b. The discrepancy between the \( N(C II) \) and \( N(C II') \) originates from the well-known problem when on the flat part of the curve of growth, that very small changes in the Doppler width can produce very large changes in the column density. Despite the complication of using the ISM average depletions for carbon and sulfur, estimating the C II column density is likely more accurately achieved by using S II.

3.5. Mg I and Mg II

Spectra of 3 (\( \eta \) UMa, \( \alpha \) Gru, and G191-B2B) of the 13 sight lines with C II' absorption also contain the Mg I and Mg II lines. We fit the optically thin Mg I line separately and the two marginally saturated Mg II lines simultaneously. The spectra and fits are shown in Figure 4, and the fit parameters are listed in Table 5. Due to the strength of the Mg II lines and relative weakness of the Mg I, not all Mg II components are detected in Mg I, although those that have consistent velocities and Doppler widths, indicating that both ions are likely part of the same collections of gas.

3.6. Individual Sight Lines

\( \alpha \) Aur.—The line of sight toward \( \alpha \) Aur is well known to exhibit a strong single interstellar absorption component (Linsky et al. 1993, 1995) that has been identified with the Local Interstellar Cloud (LIC; Lallement et al. 1995). These characteristics, complemented by the star’s proximity, contributed to it being the first LISM sight line with an \( n_e \) measurement based on absorption detected in C II and C II' (Wood & Linsky 1997). The first fit plotted in Figure 1 is for \( \alpha \) Aur. As noted by Wood & Linsky (1997), who analyzed the same data, the low signal-to-noise ratio and the saturation of the resonance line make the measurement of a precise Doppler width difficult. For this reason, we use the independently determined LISM temperature and turbulent velocity along this line of sight to force the Doppler width to be 3.48 ± 0.15 km s\(^{-1}\) (Redfield & Linsky 2004b). Our measured central velocity of \( v = 20.78 ± 0.28 \) km s\(^{-1}\) agrees with the 20.2 ± 1.5 km s\(^{-1}\) measured by Wood & Linsky (1997). Also,
we measure column densities of \( \log N(C \ion{ii}) = 14.67^{+0.14}_{-0.22} \) and \( \log N(C \ion{iii}) = 12.62 \pm 0.07 \), which agree with their \( \log N(C \ion{ii}) = 14.8 \pm 0.3 \) and \( \log N(C \ion{iii}) = 12.64 \pm 0.07 \), respectively.

\( \eta \) UMa.—The \( \eta \) UMa sight line shows two components. The Doppler width of the absorption observed toward this target is constrained based on independent temperature and turbulent velocity information (Redfield & Linsky 2004b). For the first component, \( b \) is fixed at 3.76^{+0.11}_{-0.11} \) km s\(^{-1}\), and for the second component, \( b \) is fixed at 5.60^{+0.09}_{-0.11} \) km s\(^{-1}\). All C \( \ion{n} \) fit parameters from Redfield & Linsky (2004a) agree with our measurements.
Both components are detected in Mg $\text{I}$ and Mg $\text{II}$, although the redward component is very weak and has a poorly determined central velocity and Doppler width. Indeed, the blueward (roughly $-2 \text{ km s}^{-1}$) component is consistent in velocity for both carbon and magnesium, but the measured velocity of the redward component is significantly discrepant between the two, although the errors are large. Frisch et al. (2006) provided an in-depth analysis of this particular sight line, including a measurement of $n_e \sim 0.1 \text{ cm}^{-3}$.

$\alpha$ Gru.—The spectra of $\alpha$ Gru were taken at the lowest resolution allowed in our survey. Based on this observation alone, it is difficult to identify the appropriate number of absorption components, but on the merit of past high spectral resolution
observations of other ions (e.g., Mg II, Fe II), three interstellar components have been determined along this line of sight (Redfield & Linsky 2002). We fit three components accordingly and use the same difference in central velocities between the three components as in Redfield & Linsky (2002) to constrain our C II fit. We use the wavelength solution of Fe II in order to calibrate the C II spectrum, which had no wavelength calibration image at the same wavelength. The errors are large for Doppler width and column density due to the blending of the three components. Our measurements are consistent with, although more reliable than, those by Redfield & Linsky (2004a), in which the C II resonance line was fitted alone. Our Mg II measurements agree well with those of Redfield & Linsky (2002).

**WD 0050–332.**—Using low-resolution \((R \sim 10,000; \Delta \nu \sim 30 \text{ km s}^{-1})\) *International Ultraviolet Explorer (IUE)* observations, Holberg et al. (1998) find interstellar C II toward WD 0050–332 at 10.69 ± 3.20 \text{ km s}^{-1}. At higher resolution, our C II and S II observations resolve this feature into two components at \(\sim 6.3\) and \(\sim 12.6 \text{ km s}^{-1}\). In agreement with our claim of two components, two components have also been observed in N I and Si II by Oliveira et al. (2005) at \(\sim 6.8\) and \(\sim 16.8 \text{ km s}^{-1}\). In C II, however, we observe only the stronger \(\sim 6.3 \text{ km s}^{-1}\) component. Tentative detections of highly ionized components associated with a circumstellar shell have been observed toward this target (Holberg et al. 1998; Bannister et al. 2003).

**EX Hya and IX Vel.**—Both targets have been scrutinized for observational evidence for circumbinary disks (Belle et al. 2004). In this search, which used the same data sets employed here, they identify absorption in C II, C II*, and S II, although the absorption was not identified with circumbinary absorption for either star. Instead, Belle et al. (2004) attributed the observed absorption to the LISM and fitted the profile with a single-Gaussian function, although particularly in the case of IX Vel, two components are clearly present. We fit both EX Hya and IX Vel with two components, motivated by clear asymmetries in the C II* and S II absorption. The column-weighted velocities of our components match fairly well with their estimates for EX Hya. However, the weaker C II* transition \((\lambda 1335.6627)\) appears to be used as the wavelength standard instead of the stronger transition \((\lambda 1335.7077)\), which overestimates the velocity of absorption by \(\sim 10 \text{ km s}^{-1}\). This correction improves the agreement of their C II* velocity with the other lines they measured and with our measurements. Their IX Vel estimates are significantly different from ours and

### Table 4

**Fit Parameters for S II ISM Velocity Components**

| Star Name          | Other Name | Component Number | v (km s\(^{-1}\)) | b (km s\(^{-1}\)) | log(N/CII) cm\(^{-2}\) |
|--------------------|------------|------------------|------------------|------------------|-------------------------|
| WD 0050–332        | GD 659     | 1                | 7.00 ± 0.90      | 3.07 ± 0.67      | 13.726 ± 0.065          |
|                    | GD 659     | 2                | 15.0 ± 2.9       | 2.8 ± 2.5        | 12.95 ± 0.33            |
| EX Hya             |            | 1                | −12.7 ± 2.1      | 7.4 ± 2.9        | 14.15 ± 0.12            |
|                    |            | 2                | −4.1 ± 4.0       | 2.6 ± 1.9        | 13.71 ± 0.29            |
| 74 Psc B           | HR 311     | 1                | −5.59 ± 0.93     | 3.48 ± 0.18      | 14.362 ± 0.017          |
|                    | HR 311     | 2                | 11.8 ± 1.4       | 3.91 ± 0.65      | 13.662 ± 0.052          |
| WD 0501+527        | G191-B2B   | 1                | 8.4 ± 1.2        | 2.21 ± 0.71      | 13.263 ± 0.087          |
|                    | G191-B2B   | 2                | 20.0 ± 1.6       | 4.46 ± 0.92      | 13.468 ± 0.052          |
| 74 Psc A           | HR 310     | 1                | −6.21 ± 0.92     | 4.19 ± 0.31      | 14.317 ± 0.038          |
|                    | HR 310     | 2                | 12.3 ± 2.5       | 4.6 ± 1.2        | 13.68 ± 0.11            |
| WD 0232+035        | Feige 24   | 1                | 4.11 ± 0.64      | 4.18 ± 0.95      | 13.795 ± 0.072          |
|                    | Feige 24   | 2                | 16.8 ± 3.8       | 2.6 ± 2.3        | 12.73 ± 0.38            |
| WD 2309+105        | GD 246     | 1                | −8.3 ± 1.0       | 2.42 ± 0.21      | 14.239 ± 0.035          |
|                    | GD 246     | 2                | 1.6 ± 1.5        | 4.92 ± 2.2       | 13.634 ± 0.099          |
| WD 1210+533        |            | 1                | −12.0 ± 3.6      | 5.0 ± 2.6        | 13.55 ± 0.34            |
|                    |            | 2                | −3.10 ± 0.94     | 3.0 ± 1.0        | 14.19 ± 0.12            |
|                    |            | 3                | −9.1 ± 2.4       | 4.5 ± 2.0        | 13.93 ± 0.31            |
| ρ Lup              | HR 5453    | 1                | −16.11 ± 0.79    | 5.22 ± 0.66      | 14.56 ± 0.10            |
|                    | HR 5453    | 2                | −5.3 ± 1.2       | 3.8 ± 1.2        | 13.618 ± 0.14           |
| IX Vel             |            | 1                | 20.37 ± 0.53     | 4.64 ± 0.41      | 14.320 ± 0.024          |

* Resonance-line parameters derived from simultaneous fits.

![Fig. 3](image-url) — Estimated C II column density using S II as a proxy \([\text{C II } \frac{N}{C} \text{ vs. S II}]\) vs. C II resonance-line column density derived from the saturated lines directly \([\text{C II } \frac{N}{C} \text{ vs. S II}]\). The systematic errors due to the conversion from \([\text{S II } \frac{N}{C} \text{ (solid red lines) extend beyond the random S II fitting errors. Unity (dashed line) bisects the plot window.}]

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are likely due to the poor approximation of a single Gaussian to the absorption profile. Linnell et al. (2007) also note the presence of interstellar absorption in C ii and S ii in the spectrum of IX Vel.

**G191-B2B.**—This is the third target in which the Doppler width can be constrained from independent measurements of temperature and turbulent velocity (Redfield & Linsky 2004b). For the first and second components, \( b = 4.1 \times 10^{-29} \) and \( 3.43 \times 10^{-29} \) km s\(^{-1}\), respectively. The column density measured by Redfield & Linsky (2004a) for the \( \sim 17 \) km s\(^{-1}\) component, \( \log N(C \text{ ii}) = 15.70 \pm 0.36 \), is consistent with our measurement, \( \log N(C \text{ ii}) = 15.42 \pm 0.17 \). The \( \sim 6 \) km s\(^{-1}\) component is in significant disagreement, with their estimate at \( \log N(C \text{ ii}) = 16.12 \pm 0.33 \) and ours at \( \log N(C \text{ ii}) = 14.8 \pm 0.1 \). The measurements by Redfield & Linsky (2004a) were obtained from fitting the saturated C ii resonance line only, which typically overestimates the column density. The current fits are more reliable since both C ii and C iii profiles are used simultaneously, which more tightly constrain the fit parameters. We measure \( \log N(C \text{ iii}) = 12.28 \pm 0.09 \) for the second component, which matches a velocity believed to be attributed to the LIC (Lallement et al. 1995) in this direction. Lying angularly close \((\sim 2.5^\circ)\) and believed to traverse the LIC is the line of sight toward \( \alpha \) Aur. We measure its column density to be slightly greater with \( \log N(C \text{ iii}) = 12.62 \pm 0.07 \). Since the distance to \( \alpha \) Aur is 20\% that to G191-B2B, it is clear that the LIC does not extend beyond \( \sim 13 \) pc in that direction. Mg i, Mg ii, and S ii are also measured along this line of sight using two components. The velocities agree between magnesium and sulfur, although they are slightly larger than those measured with carbon. Indeed, these C ii and C iii fits differ in velocity even from the C ii fits alone from Redfield & Linsky (2004a). Again, since the current

\[ \text{Velocity} \ (\text{km s}^{-1}) \]

**FIG. 4.—** Similar to Fig. 1, but for fits to interstellar Mg ii and Mg iii absorption toward 3 of the 13 targets that show C iii absorption. The ratios of ionization stages are used to estimate the electron density for each absorption component. Both Mg ii lines are fitted simultaneously. The parameters for these fits are listed in Table 5.
measurements are also constrained by the optically thin excited line, these measurements are more reliable. Why the G191-B2B absolute radial velocities disagree among different ions is unknown, but may be explained by systematic wavelength calibration issues, as the difference in radial velocity between the components agrees for all ions, and the spectra used for different ions are co-additions of different individual observations at various grating settings. Lemoine et al. (1996) also measure LISM absorption for two dominant components at \( \sim 9.9 \) and \( \sim 20.6 \) km s\(^{-1}\) and derive roughly similar column densities for Mg \( \text{ii} \) and C \( \text{ii} \).

\textit{WD 0232+035}.—Dupree & Raymond (1982) identified interstellar absorption in O \( \text{i} \), Si \( \text{ii} \), and N \( \text{i} \) while suggesting the presence of a Strömgren sphere due to absorption observed in C \( \text{iv} \). There have also been continued discussions of high-ionization circumstellar features for this star (e.g., Vennes & Thorstensen 1994; Holberg et al. 1998; Bannister et al. 2003). Holberg et al. (1998) observe LISM absorption of C \( \text{ii} \) at \( \nu = -4.28 \pm 2.78 \) km s\(^{-1}\) with low-resolution IUE spectra. Using the same data sets as Vennes et al. (2000), we resolve two components in both C \( \text{ii} \) and S \( \text{ii} \) at \( \sim 3.9 \) and \( \sim 17.0 \) km s\(^{-1}\), which match their \( \sim 3.1 \) and \( \sim 17.6 \) km s\(^{-1}\) components very well. We find evidence of the \( \sim 17.0 \) km s\(^{-1}\) component in S \( \text{ii} \), most notably in the strong 1259.5 Å line, whereas Vennes et al. (2000) only fit the \( \sim 3.9 \) km s\(^{-1}\) component. We observe, just as they do, only the \( \sim 3.9 \) km s\(^{-1}\) component in the C \( \text{ii} \) excited line.

\textit{WD 2309+105}.—Holberg et al. (1998) use low-resolution IUE spectra to identify one LISM component in C \( \text{ii} \) at \( \nu = -8.27 \pm 3.19 \) km s\(^{-1}\) and in all three S \( \text{ii} \) lines at velocities ranging from \(-14.5 \) to \(-6.7 \) km s\(^{-1}\). Two components are seen in high-resolution HST spectra analyzed by Oliveira et al. (2003) in N \( \text{i} \), S \( \text{ii} \), and Si \( \text{ii} \), while only the redward component is detected in C \( \text{ii} \). Using the same data set, we concur and derive nearly identical fit parameters. Our central velocities are slightly different; ours are measured at \( \sim -8.5 \pm 1.2 \) and \( 1.3 \pm 1.8 \) km s\(^{-1}\) and theirs at \( \sim -9.8 \) and \( -0.6 \) km s\(^{-1}\), but the differences are minimal given the large errors and almost identical differential velocity. Only the \( \sim -8.5 \) km s\(^{-1}\) component is detected in C \( \text{ii} \) in agreement with the IUE detection by Holberg et al. (1998).

\textit{WD 1210+533}.—Although only one component could be discerned in the low-resolution IUE spectra, Holberg et al. (1998) find C \( \text{ii} \) absorption at \( \nu = -7.32 \pm 3.21 \) km s\(^{-1}\) and C \( \text{ii} \) absorption at \( \nu = -12.14 \pm 5.92 \) km s\(^{-1}\). We detect three components in C \( \text{ii} \), while only the two redward components are seen in S \( \text{ii} \), at velocities of \( \sim -23.4 \), \(-9.3 \), and \(-3.0 \) km s\(^{-1}\). The S \( \text{ii} \) 1253.8 Å line suffers from considerable contamination, although the two components are well characterized in the other two S \( \text{ii} \) lines.

\( \rho \) Lup.—The two components characterizing the excited line of \( \rho \) Lup are similar in Doppler width and column density, lending difficulty to distinguishing the two despite a \( \sim 7 \) km s\(^{-1}\) separation. Fortunately, the asymmetric S \( \text{ii} \) absorption profile allows us to discern two LISM components. We use velocities determined from these measurements to constrain our C \( \text{ii} \) fits. Welsh & Lallement (2005) observe two components in Na \( \text{i} \) (\( \lambda 5890.0 \)) and three components in Fe \( \text{ii} \) (\( \lambda \lambda 1608.5 \), \( \lambda 21253 \), and \( \lambda 1259 \)), and Al \( \text{ii} \) (\( \lambda 1670.8 \)). Although we only find evidence for two components in C \( \text{ii} \) and S \( \text{ii} \), our components agree fairly well, with velocities \( \sim -16.1 \) and \(-9.1 \) km s\(^{-1}\). Welsh & Lallement (2005) present a thorough discussion of cloud distances and identify the \( \sim -9.1 \) km s\(^{-1}\) absorption with material at \( \sim 90 \) pc while identifying the \( \sim -16.1 \) km s\(^{-1}\) absorption with material at the neutral boundary of the Local Bubble in the direction toward \( \rho \) Lup.

4. ELECTRON DENSITY

4.1. Estimation Based on C \( \text{ii} \)

Table 6 lists the measured electron densities, \( n_e \), along the lines of sight toward the 13 targets analyzed. The values listed in the fourth column of this table are calculated using a method that compares the column densities of the resonance and excited lines of C \( \text{ii} \). Our use of this method is similar to that implemented by Spitzer & Fitzpatrick (1993) and Oliveira et al. (2003) but most parallels that of Wood & Linsky (1997).

The C \( \text{ii} \) resonance absorption line at 1334.5323 Å corresponds to the transition from the ground state \( (J = \frac{3}{2}) \), while the C \( \text{ii} ^{+} \) excited absorption lines at 1335.6627 and 1335.7077 Å correspond to the transition from the excited state of the fine-structure doublet \( (J = 3/2) \). Collisions with electrons are responsible for populating the excited state, and hence the ratio of the column densities of the two lines is proportional to the electron density. For a detailed discussion of how the fine structure of absorption lines can be used to determine density, see Bahcall & Wolf (1968).

The following relation is derived from thermal equilibrium between collisional excitation of the \( J = 3/2 \) state and radiative de-excitation:

\[
\frac{N(\text{C} ^{+} \text{ii})}{N(\text{C} ^{0} \text{ii})} = \frac{n_e C_{12}(T)}{A_{21}}.
\]

The effect of collisional de-excitation at these densities and temperatures is negligible and therefore not included in the equation above. \( N(\text{C} ^{+} \text{ii}) \) and \( N(\text{C} ^{0} \text{ii}) \) are the column densities of the resonance and excited lines, respectively. The calculation of the electron density using S \( \text{ii} \) as a proxy for C \( \text{ii} \) simply replaces \( N(\text{C} ^{0} \text{ii}) \) with \( N(\text{C} ^{+} \text{ii}) \), as derived from equation (2). The electron densities based on \( N(\text{C} ^{+} \text{ii}) \) are also listed in Table 6 (fifth column). The radiative de-excitation rate coefficient is \( A_{21} = 2.29 \times 10^{-6} \) s\(^{-1}\), as listed by Nussbaumer & Storey (1981). The collision rate coefficient can be expressed in cgs units as

\[
C_{12}(T) = \frac{8.63 \times 10^{-6} \Omega_{12}}{g_1 T^{5.5}} \exp\left(-\frac{E_{12}}{kT}\right),
\]

where the statistical weight of the ground state \( g_1 = 2 \) and the energy of the transition \( E_{12} = 1.31 \times 10^{-14} \) ergs. Since the collision strength \( \Omega_{12} \) has a very weak temperature dependence, for all targets we let \( \Omega_{12} = 2.81 \) (Hayes & Nussbaumer 1984), calculated at a temperature of 7000 K, very similar to the LISM average of 6680 K (Redfield & Linsky 2004b).

4.2. Estimation Based on Mg \( \text{i} \) and Mg \( \text{ii} \)

Three of our targets have both Mg \( \text{i} \) and Mg \( \text{ii} \) LISM absorption measurements (see Table 5). We estimate the electron density by following the procedure detailed by Lallement & Ferlet (1997) and Frisch (1994). Assuming equilibrium, the electron density can be estimated by formulating the balance of ionization and recombination between neutral and singly ionized magnesium:

\[
\frac{N(\text{Mg} ^{+} \text{ii})}{N(\text{Mg} ^{0} \text{ii})} = \Gamma + n_e \sigma_{\text{ex}},
\]

where \( \Gamma \) is the photoionization rate of Mg \( \text{i} \), \( \sigma_{\text{ex}} \) is the formation rate of Mg \( \text{ii} \) based on charge exchange, \( \alpha \) is the total recombination rate, and we have assumed that \( N(\text{Mg} ^{+} \text{ii})/n(\text{Mg} ^{+} \text{ii}) = n(\text{Mg} ^{0} \text{ii})/n(\text{Mg} ^{0} \text{ii}) \) and \( n(\text{H} ^{+} \text{ii}) = n_e \). We use a photoionization rate
### Table 6
Electron Densities

| Star Name | Other Name | Component Number | ne (cm⁻³) | n_e(C ii)⁺ | n_e(Mg ii/Mg i)⁺ | Cloud⁵ |
|-----------|------------|------------------|----------|-----------|-----------------|--------|
| α Aur     | Capella    | 1                | 0.140+0.060⁻0.039 | ...       | ...              | LIC    |
| η UMa     | Alcaid     | 1                | 0.165+0.067⁻0.039 | ...       | 0.089+0.118⁻0.087 | NGP    |
|           |            | 2                | 0.154+0.029⁻0.027 | ...       | 0.085+0.217⁻0.085 | ...    |
| α Gru     | Alnair     | 1                | 0.081+0.033⁻0.036 | ...       | <0.21            | ...    |
|           |            | 2                | 0.25+0.033⁻0.15   | ...       | 0.25+0.015⁻0.024 | (Mic, Vel) |
|           |            | 3                | 0.5+0.05⁻0.01     | ...       | 0.1+0.6⁻0.01    | LIC    |
| WD 0050–332 | GD 659  | 1                | 0.060+0.060⁻0.033 | 0.12+0.030⁻0.094 | ...       | LIC, (Cet) |
|           |            | 2                | <0.0094           | ...       | <0.11            | (Vel)  |
| EX Hya    |            | 1                | 0.004+0.003⁻0.005 | 0.094+0.034⁻0.074 | ...       | NGP, (Loo, G) |
|           |            | 2                | 0.17+0.03⁻0.01    | 0.067+0.004⁻0.041 | ...       | Gem, (Leo, Aur) |
| 74 Psc B  | HR 311     | 1                | 0.017+0.007⁻0.005 | 0.068+0.021⁻0.010 | ...       | LIC, (Hyades, Eri) |
|           | HR 311     | 2                | 0.03+0.04⁻0.01    | 0.19+0.031⁻0.012 | ...       | LIC     |
| WD 0501+527 | G191-B2B  | 1                | 0.30+0.28⁻0.28    | 0.80+0.16⁻0.15 | 0.48+0.70⁻0.48 | Hyades |
|           | G191-B2B   | 2                | 0.011+0.004⁻0.004 | 0.081+0.021⁻0.015 | <0.39     | LIC     |
| 74 Psc A  | HR 310     | 1                | 0.007+0.003⁻0.002 | 0.073+0.009⁻0.004 | ...       | ...     |
|           | HR 310     | 2                | 0.13+0.09⁻0.06    | 0.25+0.06⁻0.15 | ...       | LIC, (Hyades, Eri) |
| WD 0232+035 | Feige 23 | 1                | 0.10+0.06⁻0.04    | 0.21+0.08⁻0.04 | ...       | ...     |
|           | Feige 24   | 2                | <0.0011           | <0.15      | ...       | LIC, (G, Blue, Hyades) |
| WD 2309+105 | GD 246   | 1                | 0.12+0.012⁻0.006  | 0.084+0.010⁻0.003 | ...       | ...     |
|           | GD 246     | 2                | <0.21             | <0.060     | ...       | LIC, (Eri) |
| WD 1210+533 | ...      | 1                | <0.62             | <0.079     | ...       | ...     |
|           |            | 2                | 0.050+0.09⁻0.056  | 0.37+0.047⁻0.032 | 0.20+0.32⁻0.22 | ...     |
|           |            | 3                | 0.07+0.09⁻0.02    | 0.10+0.010⁻0.012 | ...       | LIC     |
| ρ Lup     | HR 4353    | 1                | 0.046+0.033⁻0.013 | 0.067+0.037⁻0.057 | ...       | (Gem)   |
|           | HR 4353    | 2                | 0.14+0.04⁻0.01    | 0.18+0.027⁻0.011 | ...       | (Gem)   |
| IX Vel    |            | 1                | 0.28+0.12⁻0.14    | 0.14+0.06⁻0.06 | ...       | (G, Blue) |
|           |            | 2                | 0.018+0.005⁻0.007 | 0.17+0.01⁻0.07 | ...       | (Vel)   |

⁵ Two errors listed: the first are based on the propagation of the column density errors only, while the second include errors in the cosmic abundances and the natural range of depletions of carbon and sulfur in the ISM.

Assume LISM temperature appropriate for the line of sight based on multi-ion line widths or the LISM average (Redfield & Linsky 2004b). For the second component toward η UMa, for which Redfield & Linsky (2004b) estimate a temperature of 0.44⁺0.00⁻0.00 K, we use T = 100 K.

In agreement with projected velocity and spatial distribution (Redfield & Linsky 2008). Used to calculate the weighted mean value for the LIC, n_e(LIC) = 0.12 ± 0.04 cm⁻³.

5. DISCUSSION

Figures 5–7 summarize the results of our electron density measurements. Figure 5 shows the measured electron density as a function of the C ii column density. Electron density estimates using both the saturated C ii resonance line and S ii as a proxy for C ii are shown. The measurements using S ii are more precise, even when we include systematic errors of the variation of depletion, than the C ii measurements since S ii is optically thin. Figure 6 displays the same distribution of data points, except instead of electron density, we plot the resonance-line column density versus the excited-line column density. Overplotted are constant-density contours for the mean temperature of the LISM. Both plots show that the LISM electron density measurements are relatively tightly lognormal distributed about the unweighted mean value of ~0.1 cm⁻³. Figure 7 further emphasizes the tight distribution of electron density measurements about the mean.

In particular, by using optically thin S ii lines as a proxy for C ii, we avoid systematically high column density measurements commonly derived from saturated line profiles, which ultimately lead to very low electron density determinations. For this reason, in the electron density distribution derived from S ii, we lose the low electron density tail evident with the C ii measurements, using S ii as a proxy of C ii and assuming a temperature of 8000 K. Our measurement using the same technique, assuming the LISM average temperature of 6680 K, matches quite well, n_e = 0.084+0.055⁻0.033 cm⁻³.

### 4.3 Previous Measurements

Electron density estimates have been made by other researchers for a few of the sight lines. Toward α Aur, our density (n_e = 0.140⁺0.060⁻0.039 cm⁻³) agrees well with the measured n_e = 0.11⁺0.12⁻0.05 cm⁻³ by Wood & Linsky (1997). An interstellar temperature of 6700 K (Redfield & Linsky 2004b) is included in our determination of the electron density. Vennes et al. (2000) obtained an electron density estimate toward WD 0232+035 by assuming a wide range of possible b-values that resulted in a density of n_e = 0.36 cm⁻³ for the interstellar component near 3.8 km s⁻¹. This is slightly higher, but within 2σ of our estimates, n_e = 0.10⁺0.11⁻0.08 cm⁻³ and n_e(C ii) = 0.10⁺0.14⁻0.08 cm⁻³. Oliveira et al. (2003) provide an estimate of n_e = 0.1 ± 0.01 cm⁻³ for the −9.7 km s⁻¹ component toward WD 2309+105, based on...
and the LISM average tightens to an unweighted mean of $0.11^{+0.10}_{-0.05} \, \text{cm}^{-3}$, where the 1σ errors are the dispersion about the mean value. All electron densities derived using S\,ii as a proxy range from 0.07 to 0.80 cm$^{-3}$. The distribution of $n_e$ (S\,ii) measurements matches well the $n_e$ (C\,ii) distribution, but it is significantly tighter, since the gross overestimates of the C\,ii column density are avoided. This is a strong endorsement for using S\,ii as a proxy for C\,ii, in order to make more accurate electron density measurements.

Figure 8 shows the solutions of electron density as a function of temperature for the C\,ii/C\,i technique, which does not require ionization equilibrium and has a very modest temperature dependence, and for the Mg\,ii/Mg\,i technique, which does require ionization equilibrium and has a very strong temperature dependence. This type of comparison was used by Gry & Jenkins (2001) to measure the properties along the sight line toward ε CMa. Also shown in Figure 8, for several of the sight lines, is an independent measurement of the temperature of the gas based on the widths of absorption lines of different atomic mass (Redfield & Linsky 2004b). Ideally, all three measurements should converge at the same temperature and provide a single estimate of the electron density.

5.1. LISM Cloud Properties

The last column of Table 6 indicates the interstellar cloud(s) by which the observed components may be identified. The criteria for identification include agreement with the predicted projected velocity and spatial distribution of a given LISM cloud (Redfield & Linsky 2008). Components for some of our targets (α Aur, α Gru, and G191-B2B for the LIC; η UMa for the North Galactic Pole (NGP) Cloud; G191-B2B for the Hyades Cloud) had been previously identified and assigned to specific clouds by Redfield & Linsky (2008). Many of the remaining components could also be identified with known LISM clouds (i.e., the observed radial velocity of absorption is within 3σ of the predicted projected velocity, and the line of sight traverses the spatial distribution of the cloud). Those clouds listed in parentheses also agree in velocity, but the sight line only passes near (within 20° of) the given cloud boundary. Given the similarities of the LISM cloud velocities, it is not always easy to uniquely identify cloud membership, although of the 23 C\,ii components, 11 can be firmly identified with a specific LISM cloud.

The LIC is the collection of gas that dominates LISM absorption line observations because it is detected in a large fraction of the sky. The Sun is currently located just outside of the LIC (Redfield & Linsky 2000, 2008). While LIC material has likely surrounded the Sun for the last ~100,000 yr, given the relative motion and location of the LIC and the Sun, the solar system has moved, or will very shortly move, into a different interstellar environment. However, the proximity of the LIC means that many LISM sight lines will probe this material. Indeed, 7 of
the 11 identified cloud members are of the LIC. The LIC $n_e$ measurements are shown by the shaded histogram in Figure 7, where all sight lines give a similar measurement. The $n_e (C \ ns \ s)$ measurements are used to calculate the mean when available, while $n_e (C \ ii \ a)$ is used for $\alpha$ Aur and $n_e (Mg \ ii \ / \ Mg \ i)$ is used for $\alpha$ Gru. The weighted mean is $n_e (LIC) = 0.12 \pm 0.04 \ cm^{-3}$. Dramatic variation in electron density within the LIC is not observed. However, more sight lines would be required to do a more involved investigation of intracloud variability.

Other electron density components are assigned to various other clouds, but none have as many measurements as the LIC. The NGP Cloud has two $n_e$ measurements that agree quite well with each other at $\approx 0.09 \ cm^{-3}$ and are not too different than the LIC measurement. Likewise, the Gem Cloud is observed in the second component of EX Hya ($n_e \approx 0.15$) and is similar to the LIC measurement. This is the first strong evidence that LISM clouds that are dynamically distinct from the LIC nonetheless have a similar electron density. However, not all LISM clouds share similar electron density properties. The Hyades Cloud measurement observed in the first component of G191-B2B is consistently measured at a high electron density ($\approx 0.5$) in all three techniques. The Hyades Cloud appears to be a decelerated cloud at the leading edge of the platoon of LISM clouds beyond the LIC (Redfield & Linsky 2001, 2008). It is therefore closer to G191-B2B than the LIC and without any obvious LISM clouds between it and G191-B2B to shield the ionizing radiation. The Hyades Cloud appears to have an enhanced electron density due to the photoionization of G191-B2B, while also shielding other LISM clouds, such as the LIC, from the strong ionizing radiation of G191-B2B. No obvious measurements of the second largest LISM cloud and likely future interstellar environment of the Sun, the G cloud, exist in this sample.

5.2. Sources of Ionization and Self-Shielding

Almost all of our background stars are WDs or early-type stars and therefore strong UV photon sources and significant contributors to the local radiation field. Indeed, Vallerga (1998) provided an inventory of 54 strong EUV stars that largely determine the local radiation field. The local EUV field is dominated by the B star $\epsilon$ CMa and by three WDs (Feige 24, HZ 43, and G191-B2B). Both Feige 24 and G191-B2B are identified here as having detectable C $ii$, while $\epsilon$ CMa is a relatively distant star (130 pc) and not included in our sample search, and HZ 43 did not have any detectable C $ii$, although this may be due to the significantly lower amount of LISM material in that direction. Three more of our 13 sight lines ($\alpha$ Aur, WD 0050−332, and WD 2309+105) are also included in the radiation field inventory by Vallerga (1998), and we expect that the other WDs and early-type stars that make up our remaining sight lines are also significant contributors to the local ionizing radiation field.

In the left panel of Figure 9, the electron density measurements are shown as a function of the angle from $\epsilon$ CMa. This is a test that was suggested in Vallerga (1998). Unfortunately, most sight lines cluster around 90° from the direction of $\epsilon$ CMa, but nonetheless, no clear correlation is detected. This contrasts with a possible ionization gradient toward $\epsilon$ CMa, based on hydrogen and helium column densities measured in the EUV toward many nearby WD stars (Wolff et al. 1999). Our measurements indicate that $\epsilon$ CMa does not singularly dominate the ionization structure of the LISM, and other contributors, such as the other nearby WDs and early-type stars, may be significant sources of the radiation field that dictates the electron density in the LISM.

The right panel of Figure 9 demonstrates a possible test of self-shielding in LISM clouds. The figure plots the difference of electron densities along the same sight line as a function of the total column density along the line of sight. Those with low columns should show evidence of less shielding, while large columns will provide significant shielding and a more dramatic difference in electron density measurements. No such correlation is seen, indicating that for the entire LISM sample, there does not appear to be pervasive shielding of ionizing radiation along each specific line of sight. This diagnostic assumes that for each sight line, the dominant ionization source is the background star itself, which will not necessarily be the case. A more sophisticated three-dimensional morphological model of the LISM that includes the effects of known ionizing sources is needed to accurately determine if the observed electron density is consistent with shielding of LISM clouds. However, isolated sight lines of strong ionization sources may be used to investigate the impact of shielding on electron density measurements.

As mentioned in the previous section, the G191-B2B sight line provides a possible example of self-shielding. Because many of our background targets are significant EUV sources, clouds closer to the source (i.e., more distant from the Sun) should shield the clouds farther from the source (i.e., nearer to the Sun). The LIC components offer excellent examples of this, since we know that the LIC component is tracing gas closest to the Sun. The G191-B2B sight line and others, including the $\alpha$ Gru, WD 0232+035, and WD 1210+533 sight lines, are cases where the LIC component electron density is less than the other observed component, presumably probing material from a more distant cloud that lies closer to the ionizing source (assuming that it is the background star observed). Gry & Jenkins (2001) also...
Fig. 8.—Comparison of the temperature dependence in the electron density calculation using C $^{\text{ii}}$ and C $^{\text{iii}}$ (blue) vs. Mg $^{\text{i}}$ and Mg $^{\text{ii}}$ (red). This kind of plot was used by Gry & Jenkins (2001) to put limits on the temperature and electron density of clouds along the line of sight toward e CMa. However, an independent measure of the temperature of LISM clouds is available from comparisons of the line widths of ions of different atomic mass. Redfield & Linsky (2004b) derive LISM cloud temperatures using this technique, and the temperatures for specific components are shown above by the solid vertical line. The gray scale above shows the 1, 2, and 3 $\sigma$ levels of the temperature. By combining the information in line widths and ionization abundance ratios, we can more tightly constrain both the temperature and electron density of clouds in the LISM.
noted this phenomenon in looking at the electron densities toward the most dominant ionization source, $\epsilon$ CMa. They identified two absorbers with local clouds, the LIC and Blue Cloud, which had low electron densities, $n_e = 0.08–0.17$ and $0.016–0.088$ cm$^{-3}$, respectively, while the distant absorber, which was located closer to $\epsilon$ CMa, had a significantly higher electron density, $n_e = 0.18–0.28$ cm$^{-3}$. Several counterexamples, such as GD 659 and 74 Psc A and B, are also observed, where the LIC does not have a significantly lower electron density. However, these are among the weakest ionizing sources in our sample. The 74 Psc stars are inactive and relatively cool A stars, while GD 659 is the coolest ($\sim$35,000 K) of the singleton WDs in our sample, which are all $\gtrsim$50,000 K (Lajoie & Bergeron 2007; Holberg & Bergeron 2006).

5.3. LISM Pressure Measurements

The electron density is also important for its implication for pressure measurements of LISM material. This is a particularly critical issue due to the apparent disparity of the warm LISM clouds, such as those observed by C$^\text{ii}$ absorption, which have pressures $P/k = nT \sim 3000$ K cm$^{-3}$ (e.g., Redfield 2006; Jenkins 2002), and the hot, tenuous Local Bubble gas that surrounds the warm clouds, which has pressures $\sim$10,000 K cm$^{-3}$ (e.g., Snowden et al. 1990). Recent evidence of soft X-ray emission at the heliosphere due to charge exchange between the solar wind and incoming LISM appears to contribute to the soft X-ray emission that was previously only attributed to the hot Local Bubble gas (e.g., Lallement 2004). A revised inventory of soft X-ray emission may lower the temperature and/or density of the hot gas and reduce the pressure discrepancy.

The top axis of Figure 7 shows the range of measured pressures that we obtain from our electron density measurements. This calculation assumes that temperature is constant for all sight lines, as well as a simple photoionization relationship between the electron density ($n_e$) and the neutral hydrogen density ($n_{\text{H}}$). In the case of temperature, we know that this assumption is not completely valid since we see some variation about the mean LISM value ($T = 6680$ K; Redfield & Linsky 2004a, 2004b). However, since independent temperature measurements are not available for the majority of sight lines studied here, and since the dispersion about the mean temperature is not high, it is a reasonable initial assumption. We assume $n_{\text{H}} = n_e^2 \alpha(\text{H})/\Gamma(\text{H})$ (see eq. [7] in Sofia & Jenkins 1998). The balance of the recombination rate ($\alpha$) and the ionization rate ($\Gamma$) is assumed to be constant in the LISM, and we calibrate this quantity such that the LISM average electron density ($n_e$) is consistent with the LISM measurement of $n_{\text{H}} = 0.222$ cm$^{-3}$, based on in situ measurements of $n_{\text{H}} = 0.0151$ cm$^{-3}$ (Gloeckler et al. 2004) and the H$\text{I}$–He$\text{I}$ column density ratio observed toward nearby WDs (Dupuis et al. 1995). We assume $n_H = n_e$ and $n_{\text{H}} = 0.1n_{\text{H}}$. This calculation breaks down at very low densities ($n_e < 0.007$ cm$^{-3}$), where the derived hydrogen density is less than the minimum value allowed from observations of nearby stars, measured by dividing the observed hydrogen column density by the distance to the background source (Redfield & Linsky 2004a; Linsky et al. 2000). At high densities ($n_e \geq 0.4$ cm$^{-3}$), our assumption of a constant ionization rate fails severely because the resulting column density through a typical LISM cloud ($N_{\text{H}} \geq 10^{19.5}$) leads to significant shielding of ionizing radiation (Jenkins 2004). The top axis of Figure 7 is only printed for the range of densities for which our calculation is reasonable. The distribution of electron densities then translates into an unweighted mean $P/k = 3300 \pm 5500$ K cm$^{-3}$, consistent with the range of values determined for other nearby stars using excited transitions of C$\text{I}$ by Jenkins (2002).

6. CONCLUSIONS

We analyze high spectral resolution observations of LISM absorption in order to survey the electron density in nearby interstellar material. These measurements should provide important constraints on the ionization and abundance patterns of the LISM (e.g., Slavin & Frisch 2002; Sofia & Jenkins 1998; Jenkins et al. 2000b), as well as on evolutionary models of all phases of the LISM (Breitschwerdt & de Avillez 2006). A summary of our results is as follows:

1. We searched the entire HST spectroscopic database of nearby stars (<100 pc) for detections of C$^\text{ii}$. Of the ~417 total nearby sight lines, we find 13 that show C$^\text{ii}$ absorption in 23 different velocity components. The vast majority of these detections are new.

2. Using the C$^\text{ii}$–C$^\text{I}$ ratio, we infer the electron density. To increase the accuracy of our results, particularly in terms of measuring the column density of the saturated C$^\text{ii}$ resonance line, we employ three analytical strategies: (a) simultaneously...
fitting both the C ii and C iii profiles, allowing the optical thin C ii line to constrain the line width; (b) using independently derived temperatures from comparison of line widths to constrain the acceptable range of line widths for C ii; and (c) using easily measured S iii column densities as a proxy for C iii column density.

3. The distribution of electron densities based on using S iii as a proxy for C iii is similar to the distribution based on carbon alone, while significantly tighter. This is a promising technique to avoid grossly overestimating the C iii column density based on the saturated line profile.

4. We find that the distribution of measured LISM electron densities \((n_e)\) is consistent with a lognormal profile, with a mean (unweighted) value of \(n_e(C\, ii) = 0.11^{+0.10}_{-0.05} \, \text{cm}^{-3}\).

5. We assign individual velocity components to specific LISM clouds based on kinematical and spatial properties. In particular, the LIC is probed by seven different sight lines, which all give roughly identical electron density measurements. The weighted mean value for the LIC is \(n_e = 0.12 \pm 0.04 \, \text{cm}^{-3}\).

6. Two clouds, the NGP and Gem Clouds, show similar electron density properties as the LIC. The Hyades Cloud, a decelerated cloud at the leading edge of the platoon of LISM clouds, has a significantly higher electron density than the LIC. Observed toward G191-B2B, the high electron density may be caused by the lack of shielding from such a strong radiation source.

7. Almost all of our background sources are significant ionizing sources that may influence the ionization structure and thereby the electron density of the gas along the line of sight. We do not find evidence that the ionization structure of the LISM is dominated by a single source, namely, \(\epsilon\) CMa.

8. We see evidence of more distant clouds (i.e., those closest to the ionizing sources) shielding nearer clouds (i.e., those farther from the ionizing sources). In several examples, the LIC component that is known to be farthest from the ionizing source has a lower electron density than the component along the same line of sight that is nearer to the radiation source. Although counterexamples exist, they are toward the weakest radiation sources in our sample.

9. The range in electron density is used to estimate the range of pressures that may be found in warm LISM clouds. Given simple assumptions, the measured electron densities correspond to an unweighted mean pressure \(P/k = 3300^{+5500}_{-1900} \, \text{K cm}^{-3}\).

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