MULTIDIMENSIONAL SIMULATIONS FOR EARLY-PHASE SPECTRA OF ASPHERICAL HYPERNOVAE: SN 1998bw AND OFF-AXIS HYPERNOVAE

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ABSTRACT

Early-phase optical spectra of aspherical jet-like supernovae (SNe) are presented. We focus on energetic core-collapse SNe, or hypernovae. Based on hydrodynamic and nucleosynthetic models, radiative transfer in the SN atmosphere is solved with a multidimensional Monte Carlo radiative transfer code, SAMURAI. Since the luminosity is boosted in the jet direction, the temperature there is higher than in the equatorial plane by ~2000 K. This causes anisotropic ionization in the ejecta. Emergent spectra are different depending on viewing angle, reflecting both aspherical abundance distribution and anisotropic ionization. Spectra computed with an aspherical explosion model with kinetic energy \(2 \times 10^{51}\) ergs are compatible with those of the Type Ic SN 1998bw if ~10%–20% of the synthesized metals are mixed out to higher velocities. The simulations enable us to predict the properties of off-axis hypernovae. Even if an aspherical hypernova explosion is observed from the side, it should show hypernova-like spectra but with some differences in the line velocity, the width of the Fe absorptions, and the strength of the Na I line.

Subject headings: radiative transfer — supernovae: general — supernovae: individual (SN 1998bw)

1. INTRODUCTION

The connection between the long gamma-ray bursts (GRBs) and a special class of Type Ic supernovae (GRB-SNe) is now well established (Galama et al. 1998; Hjorth et al. 2003; Stanek et al. 2003; Malesani et al. 2004; Pian et al. 2006). Since GRBs are induced by relativistic jets, GRB-SNe are also expected to be aspherical. There is also increasing evidence that core-collapse SNe are not spherically symmetric, coming from, e.g., the detection of polarization in several SNe (e.g., Wang et al. 2001; Kawabata et al. 2002; Leonard et al. 2006) and late-time spectroscopy (Mazzali et al. 2005).

However, the progenitors, the explosion mechanisms, and the origins of diverse properties of GRB-SNe are still not understood (e.g., Nomoto et al. 2007). In order to answer such questions, it is crucial that the properties of the explosions (e.g., the mass and kinetic energy of the ejecta and the asymmetry of the explosion) are accurately derived from observations. For this purpose, we should know the observational properties of multidimensional SN explosions and to what degree they are affected by orientation effects.

This is an area that has not been studied in great depth, despite its importance. Although some works have addressed the properties of the light curve (LC) or the late-phase spectra of aspherical SNe (e.g., Höflich et al. 1999; Maeda et al. 2002, 2006a, 2006b; Sim 2007), only a few studies of early-phase spectra (t \(\sim\) 50 days, where \(t\) is the time since the explosion) have been performed (e.g., Höflich et al. 1996 for core-collapse SNe; Thomas et al. 2002; Kasen et al. 2004; Tanaka et al. 2006 for Type Ia SNe).

SNe that associate with GRBs are thought to be highly energetic explosions, i.e., hypernovae (here defined as SNe with ejecta kinetic energy \(E_{S} = 10^{51}\) ergs > 10; e.g., Nomoto et al. 2006) indicated by broad line features in their early-phase spectra. However, this suggestion was based on analyses that assumed spherical symmetry. No realistic multidimensional explosion models have been verified against the observed early-phase spectra.

We report the first study of early-phase spectra of aspherical hypernova models using a multidimensional radiative transfer code. The synthetic spectra are compared with observed spectra of SN 1998bw, and implications for off-axis hypernovae are discussed.

2. MODELS

We use the results of multidimensional hydrodynamic and nucleosynthetic calculations for SN 1998bw (Maeda et al. 2002) as input density and element distributions. In the present simulations, 16 elements are included, i.e., H, He, C, N, O, Ne, Mg, Si, S, Ti, Cr, Ca, Ti, Fe, Co, and Ni. Since the original models used a He star as a progenitor, we simply replace the abundance of the He layer with that of the C+O layer. In the hydrodynamic model, energy is deposited aspherically, with more energy in the jet direction (\(z\)-axis). As a result, \(^{56}\)Ni is preferentially synthesized along this direction (see Fig. 3 of Maeda et al. 2002). In this Letter, an aspherical model with \(E_{S} = 20\) (A20) and two spherical models with \(E_{S} = 20\) and 50 (F20 and F50, respectively) are studied (Table 1). They are constructed on the basis of the models with \(E_{S} = 10\) (Maeda et al. 2002, 2006a).

Bolometric LCs computed in three-dimensional space by Maeda et al. (2006a) are also used as input for the spectral calculations. A common problem in modeling hypernova LCs is that a spherical model reproducing the early rise of the LC (\(E_{S} = 50\) for SN 1998bw) declines more rapidly than the observed LC at \(t\sim 100–150\) days (Nakamura et al. 2001; Maeda et al. 2003). This problem can be solved by aspherical models with a polar view. In aspherical models, even with a lower kinetic energy (\(E_{S} = 10–20\) than in the spherical case (\(E_{S} = 50\)), which allows sufficient trapping of \(\gamma\)-rays at late times, the rapid rise of the LC can be reproduced because of the extended \(^{56}\)Ni distribution (Maeda et al. 2006a).

3. METHOD

In order to study the detailed properties of the radiation from aspherical SNe, we have unified a Supernova MUltidimensional RADiative transfer code (SAMURAI). SAMURAI is a combination
TABLE 1
SUMMARY OF EXPLOSION MODELS

| Model Name | $M_e$ | $E_{51}$ | $M^{56Ni}$ | $\rho_c$ |
|------------|-------|----------|------------|---------|
| A20 ....... | 10    | 20       | 0.39       | 0.0     |
| F20 ....... | 10    | 20       | 0.31       | 0.0     |
| F50 ....... | 10    | 50       | 0.40       | 0.0     |
| A20p0.2 .... | 10 | 20      | 0.39       | 0.2     |

- $M_e$: The mass of the ejecta ($M_e$).
- $E_{51}$: The ejected $^{56}$Ni mass ($M_e$).
- $M^{56Ni}$: The mixing fraction (see § 5).

4. RESULTS

At $t = 20$ days, the photosphere of model A20 is mildly aspherical with an axis ratio of 1.4 between the polar and equatorial directions. The emergent luminosity is also anisotropic, being brighter by a factor of 1.2 in the polar direction (Maeda et al. 2006a). In contrast, the local luminosity at the photosphere, $L_{\text{in}}(\theta)$, is highly aspherical, with an axis ratio of 6.3, tracing the aspherical $^{56}$Ni distribution. At this epoch, the photosphere is still outside the region where explosive nucleo-
The ionization fractions of Ca II of model A20 at 30 days after the explosion. The upper left panel of Figure 1 shows the computed temperature structure in the atmosphere of model A20 at 30 days after the explosion. Upper right, lower left panels: Ionization fractions of Ca II and Ca III. Lower right panel: Sobolev optical depth of Ca II λ8542.

The upper left panel of Figure 1 shows the calculated temperature structure. The boosted luminosity makes the temperature near the z-axis higher than in the equatorial plane by ~2000 K. This anisotropy of the temperature causes anisotropies of the ionization structure. The upper right and lower left panels of Figure 1 show the ionization fractions of Fe II and Fe III, respectively. Because of the high temperature, Fe II is suppressed by a factor of more than 30 near the z-axis. As a result, the optical depth of the Fe II lines is larger in the equatorial plane (Fig. 1, lower right panel).

The upper panel of Figure 2 shows the synthetic spectra at t = 20 days for models A20 (red, green, and blue: θ = 0° [polar], 45°, and 90° [equatorial], respectively), F20 (black), and F50 (gray). Here θ is measured from the z-axis. All absorption lines except for Si II λ6355 are stronger for larger θ, i.e., for an equatorial view. This is because all species that have strong lines, i.e., O i, Si ii, Ca ii, Ti ii, Cr ii, and Fe ii, dominate near the equator but not near the z-axis as shown in Figure 1 for the case of Fe. This also leads to a more effective flux-blocking in the near-UV in the side-viewed spectrum.

The lower panel of Figure 2 shows the synthetic spectra at t = 30 days. At this epoch, the local luminosity and the temperature structure are still anisotropic, although the photosphere is almost spherical (Fig. 3, upper left panel). Consequently, the distribution of ionization fractions is also anispherical (see Fig. 3, upper right and lower left panels, for the case of Ca). However, the emergent spectra are not significantly different for different viewing angles (Fig. 2). This is the effect of the aspherical abundance distribution. The photosphere at this epoch is located inside the region where heavy elements are synthesized in the explosion. Since nucleosynthesis occurs occurs entirely near the polar direction in our model, the suppression of important ions near the z-axis is compensated by the larger abundance of the heavy elements (Fig. 3, lower right panel). Only the Na i λ5890 line is stronger for larger θ because of the combined effect of the lower temperature that favors Na i and the higher abundance of Na, which is predominantly synthesized before the explosion.

Given the strong anisotropy in L⊥(θ), the side-viewed spectrum of model A20 consists mainly of photons escaping from the ejecta of θ ~ 60°. Since physical quantities there [e.g., isotropic mass and kinetic energy of the ejecta in that direction, and L⊥(θ)] are similar to those of model F20, the spectrum of model F20 is similar to the polar-viewed spectrum of model A20. Model F50 shows a very strong Na i λ5890 line owing to the low temperature.

6 Si ii λ6355 is not strong at θ = 90° because the temperature in the equatorial plane is too low to activate the transition.

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5. COMPARISON WITH SN 1998bw

The spectra computed for the aspherical model A20 are compared with those of SN 1998bw in Figure 4. Because of the detection of GRB 980425, it is tempting to compare the polar-viewed spectrum (red lines) with that of SN 1998bw. Since the computed LC is brighter than the observation by a factor of 1.25 at $t = 20$ days (Maeda et al. 2006a), the synthetic spectra at $t = 20$ days are scaled after the computation by this factor in Figure 4.

At $t = 20$ days (Fig. 4, upper panel), the scaled synthetic spectrum of model A20 with $\theta = 0^\circ$ does not show the broad absorption trough between 4000 and 5000 Å that is typical of hypernovae such as SN 1998bw. This is because in the model Fe line absorption is not effective at high velocities, making it possible for photons to escape effectively at 4000–5000 Å. Similarly, the O I–Ca II absorption at 7000–8000 Å is very weak in model A20. The Si line velocity in the model is, however, comparable to that in SN 1998bw.

At $t = 30$ days (Fig. 4, lower panel), the Ca II and Fe II lines are somewhat stronger, although the O I–Ca II absorption at 7000–8000 Å is still narrower than in the observed spectrum. The peaks around 4000 and 4500 Å are partially suppressed for a polar view by the high-velocity absorption by the extended Fe near the jet, while they are strong in the synthetic spectrum for the equatorial view. The suppression of the peaks is similarly seen in the spectrum of SN 1998bw. The Na I λ5890 line is very weak in SN 1998bw, in analogy with the spectrum with $\theta = 0^\circ$.

The strengths of the Ca II and Fe II lines at $t = 20$ days can be increased if heavy elements synthesized in the explosion are mixed to outer layers. In SN explosions, Rayleigh-Taylor (R-T) instabilities are expected to occur, which could deliver the newly synthesized elements to higher velocities. We simulate this possibility by introducing a parameter, the mixing fraction $p$. If the mass ejected inside a conical section of the ejecta is $M_\text{e}^i(\theta)$, and $M_\text{e}(\theta)$ is the mass of a certain element contained in the conical section, we simulate mixing by taking a fraction ($p$) of the mass of a newly synthesized element and distributing this mass homogeneously in the conical section. Thus, in the outer layers, the newly synthesized element will have an abundance $X_p(r, \theta) = pM_\text{e}^i(\theta)/M_\text{e}(\theta)$ (plus pre-SN abundance). In the deeper region, the abundances are redistributed by a mixing down of the unburned material.

In Figure 4, synthetic spectra of models with $p = 0.2$ (model A20p0.2) are shown. At $t = 20$ days (the synthetic spectra are similarly scaled; upper panel of Fig. 4), although the absorption line at 7000–8000 Å is still weaker than the observation, the Fe lines are strong, causing the broad absorption at $\leq 5000$ Å followed by the peak at $\sim 5200$ Å. At $t = 30$ days, the strong Ca II feature contributes to the broad absorption at 7000–8000 Å, which is comparable to that in SN 1998bw.

Although the agreement of the spectra is far from perfect, the spectra of the model with mixing are in qualitative agreement with those of SN 1998bw. Since we use multidimensional hydrodynamic and nucleosynthetic models, which are computed ab initio, there are not many parameters that we can freely control. The agreement could be improved by additional modifications, but they would be complicated and are beyond the scope of this work. Note that our input luminosity is higher than that of SN 1998bw at $t = 20$ days, making the ejecta temperature too high. If a luminosity comparable to that of SN 1998bw is used, neglecting the possible displacement of the photosphere, the O I–Ca II feature becomes as strong as the observation. Even in this case, however, mixing with $p \sim 0.1$ is still required.

6. CONCLUSIONS

We have presented the first detailed simulations of the early-phase optical spectra of realistic jet-like hypernova models. The emergent spectrum is different for different viewing angles. The spectral properties are determined by the combination of aspherical abundances and anisotropic ionization states.

Considering the complexity of the problem, the synthetic polar-viewed spectra are in reasonable agreement with those of SN 1998bw, when $\sim 10\%–20\%$ of the synthesized material is mixed out to higher velocities. The kinetic energy of an aspherical model that reproduces, at least qualitatively, the spectra of SN 1998bw ($E_{51} = 20$) can be smaller than that of a well-fitting spherical model ($E_{51} = 50$). This is consistent with previous results obtained from models of the LC and the late-time spectra (Maeda et al. 2006a, 2006b).

The simulations enable us to predict the properties of the early-phase spectra of hypernovae viewed off-axis. Compared with the spectra seen from the polar direction (red), the spectra seen from the equatorial direction (blue) have (1) a slightly lower absorption velocity, (2) stronger peaks around 4000 and 4500 Å, and (3) a stronger Na I λ5890 line than in SN 1998bw. SNe similar to SN 1998bw in ejecta mass, kinetic energy, and mass Ni should always show spectral features of “hypernova” or “broad-line supernovae” but with some differences as described above, depending on the viewing angle.

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