The Effective Quintessence from String Landscape

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Abstract

The quintessence-like potential of vacuum energy can meet requirement from quantum gravity as well as the accelerating expansion of the universe. Since string theory only admits vacuum with negative energy density, one needs to lift the anti-de Sitter vacuum to the meta-stable de Sitter vacuum with positive vacuum energy density to explain the accelerating expansion of the universe. However the two criterion about the swampland conjecture destroys the effort. Based on the possible large scale Lorentz violation, we define an effective cosmological constant which depends not only on the bare cosmological constant known as the vacuum energy density but also on the Lorentz violation effect. We find the evolution of the effective cosmological constant exhibits the behavior of quintessence potential driving universe when the bare cosmological constant is from string landscape regime while the effective cosmological constant exhibits local minimum during evolution which certainly is caused by a meta-stable de Sitter potential in the swampland regime. The critical value of of the transition for the effective cosmological constant is approximately zero. The frozen large scale Lorentz violation can uplift the AdS vacua to an effective quintessence-like one in this sense.

1 Introduction

Quantum theory of gravity is the all the time seeking goal of physics since the success of quantum field theory of electroweak and strong interactions. String theory is the most hopeful candidate of the goal. However, one encounters the puzzle that there are too many possible vacua, of the order of $10^{500}$, in string theory when compactifying the extra six dimension to get a four dimensional effective low energy theory. The vacuum energy of different cacua constitute a complex landscape\cite{1}. Anti-de Sitter type of vacua from flux compactification of string are generic and supersymmetry preserving. To account for the late time accelerating expansion of the universe, one need to lift the negative anti-de Sitter vacuum energy to a meta-stable positive de Sitter one with some unnatural techniques in string theory\cite{2, 3}. However the low effective field theory built on these meta-stable de Sitter vacuum do not have a UV completion and are not supposed to be consistent with quantum gravity. Vafa et.al. propose the idea of swampland to describe the low energy physics models which look consistent but ultimately are not when coupled with gravity\cite{11, 14}. They propose two criteria to distinguish the swampland from landscape\cite{5, 6}. The second swampland criterion forbid the meta-stable dS potential as allowed but AdS and quintessence can pass the criteria’s requirement. A refined version of swampland condition is proposed soon after the original one which turns out to evade all counter-examples at scalar potential maxima that have been raised but potential with local minimum or the meta-stable dS type is still not allowed\cite{7}. To explain the accelerating expansion by ΛCDM model one seems to need a stable or meta-stable dS vacua. Unfortunately, the tension between swampland condition and inflation excludes the dS vacua leaving possible scalar potential with allowed quintessence or AdS types\cite{6}. Does the universe with an AdS vacua expands accelerating? Is the quintessence a fundamental canonical scalar field or the quintessence type of potential emerged from other fundamental effect? We try to explore the question in a different view of point.

It is widely accepted that all the four kinds of basic interactions will be unified into quantum gravity at the Planck energy scale and the universe is dominated by quantum gravity at the beginning of creation.

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The theory describing quantum gravity is far from completion. String theory is one of the efforts. There are many different approaches, such as loop quantum gravity, Hořava-Lifshitz gravity, non-commutative geometry etc. Many of them share the common feather of explicit Lorentz violation and part of them are non-local theories and implicit Lorentz violated [8, 9].

The observable universe is believed to be expanded from a very tiny part of spacetime in the era of quantum gravity domination by inflation. The scale of area of Lorentz violation is expanded so fast as to exceed the horizon in a very short time. Parts of the region interacting via quantum gravity lose interaction at the instance by inflation and Lorentz violation of the small region before inflation is frozen in the large scale after inflation in this way. The scale may re-enter the horizon during the period of normal expansion depending on details of quantum gravity and inflation physics [10].

There are also possible signs of large scale Lorentz violation in the cosmic observations, e.g. the anisotropies in the low-l multi-pole expansion of CMB power spectrum. The normals of quadra-pole and octopole are not coincided with the direction of dipole which indicates the boost transformation from CMB static frame to the peculiar motion frame is not a simple Lorentz boost [11].

General relativity is a very successful gravitation theory at least at small scale. When it is applied to the larger scale, there appear some deviations from prediction of general relativity though the deviations are usually not recognized as deviations but as caused by unknown energy-momentum distribution known as dark matter and dark energy. One has to take dark energy into account when dealing with the cosmic scale physics. Due to the possible frozen Lorentz violation at large scale, the effective theory of gravity at the cosmic scale has to take Lorentz violation into account and it will be inevitably a non-Einstein gravitation theory for general relativity is local Lorentz invariant. Actually the ΛCDM model is very successful at explaining most cosmological phenomenon by adding cosmological constant term into Einstein theory of gravity,

\[
S_E = \frac{1}{16\pi G} \int d^4x \sqrt{-g} \left( R - 2\Lambda \right)
\]

or

\[
R_{\mu\nu} - \frac{1}{2} R g_{\mu\nu} + \Lambda g_{\mu\nu} = 8\pi G(T_{M})_{\mu\nu},
\]

where \( \Lambda \) is the cosmological constant which can also be regarded as the vacuum energy density. However, there is a big puzzle about cosmological constant when it is regarded as the vacuum energy density for theoretical prediction of its value is of 54 to 112 order of 10 higher than the value obtained by observation. One has to take fine tuning as a solution which lowers the confidence level of the model to a situation hard to accept [12].

2 Gravity with Large Scale Lorentz Violation

The idea of large scale Lorentz violation is first proposed in 2015 [13]. The framework of constructing large scale effective gravity with LV is build up with gauge principle via the equivalence principle by utilizing the constrain dynamics and very special relativity (VSR) symmetry as a modified gravity model [13–16]. Take \( SIM(2) \) symmetry, the proper subgroup of Lorentz group, as an example in the language of tetrad field \( h_{\alpha}^{\mu} \). One can build local \( SIM(2) \) gravity with gauge principle by constraining the gauge potential, the Lorentz connection, takes value only on \( sim(2) \) algebra. The action for a \( sim(2) \) gravity can be taken as

\[
S_{sim2} = \frac{1}{16\pi G} \int d^4x h \left( R^{ab}_{\alpha\beta} + \lambda_1 {A}^{10}_{\alpha\beta} - A^{31}_{\alpha\beta} + \lambda_2 {A}^{20}_{\alpha\beta} + A^{23}_{\alpha\beta} \right)
\]

where the Lagrange multipliers are used to constrain the Lorentz connection taking values on \( sim(2) \) algebra and \( h = \det h_{\alpha}^{\mu} \). The Lagrange-multipliers terms contribute an effective angular momentum distribution \( C_{Meff} \) such that

\[
D_{\nu} \left( h_a^{\nu} h_b^{\mu} - h_a^{\mu} h_b^{\nu} \right) = 16\pi G \left( C_M + C_{Meff} \right)_{ab}^{\mu
u}
\]

where \( C_M \) is the angular momentum distribution of the matter source. This equation leads to the torsion free condition in general relativity in the case of \( C_M = 0 \) where \( C_{Meff} = 0 \) always holds for the local Lorentz invariance. On the contrary, Lorentz violation leads to non-trivial distribution of torsion tensor.
even in the scalar matter source case where \( C_M = 0 \). The variation respect to tetrad fields \( \frac{\delta S_{\text{sim2}}}{\delta h^a_{\mu}} \) gives
the equations of motion for tetrad fields,

\[
R^a_b - \frac{1}{2} R \delta^a_b = \frac{8\pi G}{c^4} (T_M)^a_b ,
\]

where \( (T_M)^a_b \) is the energy-momentum tensor of matter source. It should be noted that the connection hides in \( R^a_b \) and \( R \) is no longer the torsion free Levi-Civita one for the reason discussed in [1]. Generally, the connection can be decomposed into torsion-less part and contortion

\[
A^a_{bc} = \tilde{A}^a_{bc} + K^a_{bc} ,
\]

where \( \tilde{A}^a_{bc} \) is the Levi-Civita connection and \( K^a_{bc} \) is the contortion part. The curvature can be decomposed into three parts with the help of decomposition of spin connection as

\[
R^{mn}_{\ ab} = \tilde{R}^{mn}_{\ ab} + R^K_{\ mn\ ab} + R_{CK}^{mn\ ab} ,
\]

where \( \tilde{R}^{mn}_{\ ab} \) and \( R^K_{\ mn\ ab} \) are the curvatures composed of torsion-free connection and contortion respectively, while \( R_{CK}^{mn\ ab} \) contains cross terms of them. In the geometric unit \( \frac{8\pi G}{c^4} = 1 \), we can rewrite (13) as

\[
\tilde{R}^a_b - \frac{1}{2} \delta^a_b \tilde{R} = (T_{eff} + T_M)^a_b ,
\]

where \( \tilde{R}^a_b \) and \( \tilde{R} \) are generated by torsion free Levi-Civita connection \( \tilde{A}^a_{bc} \) and \( T_{eff} \) collects all the terms involving contortion \( K^a_{bc} \).

The effective energy-momentum tensor \( T_{eff} \) contributes to the gravitation in addition to matter contribution \( T_M \) and appears as the dark partner of the matter distribution. The non-trivial effective contribution to the energy-momentum distribution by contortion is expected to be responsible for the dark partner of the matter. The Bianchi identity guarantees the conservation of \( T_{eff} \) respected to the Levi-Civita connection \( A^a_{bc} \).

Based upon our analysis on possible frozen large scale Lorentz violation, the cosmological observable actually are not Lorentz covariant but still \( SO(3) \) covariant. Simply restricting the components of Lorentz gauge field \( A^a_{\mu} \), nontrivial only on \( SO(3) \) generators in the construction of a gravitation theory with Lorentz boost violation would result in a degenerated dynamics. The reason is that boost transformation is not prohibited at the large scale actually for only the Lorentz boost transformation is violated. There are discussions on the modification of Lorentz algebra at quantum level by Hopf algebra or deformed Poincare algebra such as the \( \kappa \)-Poincare etc. as well as other quantum gravity model like Horava-Lifshitz gravity in which the Lorentz boost is automatically violated.

Observing that the Lorentz gauge potentials transform as

\[
A^a_{\ mu} = \Lambda^a_c (x) A^c_{\ mu} \Lambda_b^d (x) + \Lambda^a_c (x) \partial_{\mu} \Lambda_b^c (x)
\]

under local Lorentz transformation \( \Lambda(x) \), it is obvious that \( A_{\ i\mu} = \Lambda^j_i (x) A^0_{j\mu} \) for a rotation transformation \( \Lambda \in SO(3) \). Hence the restriction to the Lorentz gauge potentials can be proposed as

\[
(A^0_{1\mu})^2 + (A^0_{2\mu})^2 + (A^0_{3\mu})^2 = (f_\mu(x))^2
\]

where \( f_\mu(x) \) can be regarded as a measurement of the magnitude of the boost violation in some sense, and it is invariant under a local \( SO(3) \) gauge transformation on tetrad fields \( h_{\ a\mu} \) respected to the tetrad indices but frame dependent.

The action of effective gravity with large scale Lorentz violation can then be given by,

\[
S=\frac{c^4}{16\pi G} \int d^4 x h \left( R - 2\Lambda_0 + \lambda^a \left( (A^0_{1\mu})^2 + (A^0_{2\mu})^2 + (A^0_{3\mu})^2 - f_\mu^2 \right) \right)
\]
where the repeated superscript and the subscript $\mu$ of $\lambda^\nu$ and $A^{0\mu}$ respectively mean summation and $\Lambda_0$ is the bare cosmological constant given by the vacuum energy density.

The variation respect to tetrad fields $\frac{\delta S}{\delta h_{a\mu}}$ gives the equations of motion for tetrad fields,

$$G^a_{\ b} \equiv R^a_{\ b} - \frac{1}{2} R \delta^a_{\ b} + \Lambda_0 \delta^a_{\ b} = \frac{8 \pi G}{c^4} (T_{\text{M}})^a_{\ b},$$  \hspace{1cm} (13)

where $(T_{\text{M}})^a_{\ b}$ is the energy-momentum tensor of all the matter source both luminous and dark one. As discussed in eq. (5), $R^a_{\ b}$ and $R$ are composed of connection with torsion. We can rewrite eq. (13) in the form of eq. (8) with curvature composed only of torsion-less Levi-Civita connection and the contortion for the $i,j$ indices combination. With the decomposion $A^a_{\ b\mu}$ into Levi-Civita connection $\widetilde{A}^a_{\ b\mu}$ and contortion $K^a_{\ b\mu}$ in eq. (6), Eqs. (16) can be expressed in detail as

$$K^0_{\ 12} = K^0_{\ 21}, \ K^1_{\ 23} = 0, \ K^2_{\ 12} = -K^0_{\ 10}, \ K^3_{\ 13} = -K^0_{\ 10},$$
$$K^0_{\ 23} = K^0_{\ 32}, \ K^2_{\ 31} = 0, \ K^3_{\ 23} = -K^0_{\ 20}, \ K^1_{\ 21} = -K^0_{\ 20},$$
$$K^0_{\ 31} = K^0_{\ 13}, \ K^1_{\ 32} = 0, \ K^2_{\ 31} = -K^0_{\ 30}, \ K^3_{\ 23} = -K^0_{\ 30}.$$  \hspace{1cm} (17)

and eqs. (15) as

$$2K^0_{\ 10}h^0_{\ \mu} + (K^0_{\ 22} + K^0_{\ 33})h^1_{\ \mu} - (K^1_{\ 20} + K^0_{\ 21})h^2_{\ \mu} + (K^3_{\ 10} - K^0_{\ 31})h^3_{\ \mu} + \lambda^\mu (A^1_{\ 10}h^0_{\ \mu} + A^0_{\ 11}h^2_{\ \mu} + A^0_{\ 13}h^3_{\ \mu}) = 0$$
$$2K^0_{\ 20}h^0_{\ \mu} + (K^1_{\ 20} - K^0_{\ 12})h^1_{\ \mu} + (K^0_{\ 11} + K^0_{\ 33})h^2_{\ \mu} - (K^2_{\ 30} + K^0_{\ 32})h^3_{\ \mu} + \lambda^\mu (A^0_{\ 20}h^0_{\ \mu} + A^0_{\ 21}h^1_{\ \mu} + A^0_{\ 22}h^2_{\ \mu} + A^0_{\ 23}h^3_{\ \mu}) = 0$$
$$2K^0_{\ 30}h^0_{\ \mu} + (K^3_{\ 10} + K^0_{\ 13})h^1_{\ \mu} + (K^2_{\ 30} + K^0_{\ 23})h^2_{\ \mu} + (K^0_{\ 11} + K^0_{\ 22})h^3_{\ \mu} + \lambda^\mu (A^0_{\ 30}h^0_{\ \mu} + A^0_{\ 31}h^1_{\ \mu} + A^0_{\ 32}h^2_{\ \mu} + A^0_{\ 33}h^3_{\ \mu}) = 0.$$  \hspace{1cm} (18)

Actually the tetrad fields $h^a_{\ \mu}$ satisfy eqs. (17), (18) and eq. (13) as well as the constrains conditions [11] simultaneously together with the contortion $K^a_{\ b\mu}$ and the $\lambda$ multipliers. To solve the simultaneous equations, one can employ the cosmological principle which holds in the CMB static reference frame. The symmetry requirement by cosmological principle gives that the metric for the universe must have the form of Robertson-Walker space-time metric,

$$ds^2 = dt^2 - a(t)^2 \left( \frac{dr^2}{1 - kr^2} + r^2d\theta^2 + r^2 \sin^2 \theta d\phi^2 \right).$$  \hspace{1cm} (19)

The instantaneous co-moving tetrad basis can be read out directly,

$$h^0 = dt, \ h^1 = \frac{a(t)}{\sqrt{1 - kr^2}} dr, \ h^2 = a(t)r d\theta, \ h^3 = a(t)r \sin \theta d\phi.$$  \hspace{1cm} (20)
The cosmic media should be a perfect fluid by the analysis of cosmological principle. The energy momentum $T_M$ has the character of a perfect fluid and can be described by the energy density $\rho$ and pressure $p$, i.e. with the diagonal form of $(T_M)^a{}_b = \text{diag}(\rho, -p, -p, -p)$. The perfect fluid energy momentum tensor $T_M$ in eq.(13) requires $G^{\alpha \beta} = 0, \forall \alpha \neq \beta$. One can get the conclusion that the only possible non-zero independent components of contortion $K^a{}_{bc}$ are $K^0{}_{11}, K^0{}_{22}, K^0{}_{33}$ while the others can all be determined to be zero.

The cosmological principle also requires all cosmic physical quantities depend only on cosmic time $t$, and hence

$$K^0{}_{11} = K^0{}_{22} = K^0{}_{33} = \dot{\mathcal{X}}(t) \quad (21)$$

are the solutions required. The dependence between $\mathcal{X}(t)$ and $f_\mu(x)$ can be derived as

$$(f_t, f_r, f_\theta, f_\varphi) = (a(t), \dot{\mathcal{X}}(t) + \dot{a}(t)) \cdot \left(0, \frac{1}{\sqrt{1 - k^2}}, r, r \sin \theta \right) \quad (22)$$

It is a reasonable result that the only one remaining of the four degrees of freedom for $f_\mu(x)$ is expressed by $\mathcal{X}(t)$ while the other three are used to fix the reference frame $h_{\alpha \mu}$.

3 Late time expansion of the universe in landscape and swamp-land

Denote the contribution to the energy momentum tensor $(T_{eff})^a{}_c$ of eq.(14) in the co-moving frame of Robertson-Walker universe from dark partner of a perfect fluid energy momentum tensor $(T_M)^a{}_b = \text{Diag}(\rho, -p, -p, -p)$ as

$$T_K^a{}_c \equiv (T_{eff})^a{}_c = \text{Diag}(\rho_K, -p_K, -p_K, -p_K) \quad (23)$$

We can get

$$\rho_K = - \left(3\dot{\mathcal{X}}^2 + 6\dot{\mathcal{X}} \frac{\dot{a}}{a} - \Lambda_0 \right) \quad (24)$$

and

$$p_K = \dot{\mathcal{X}}^2 + 4\dot{\mathcal{X}} \frac{\dot{a}}{a} + 2\dot{\mathcal{X}} - \Lambda_0 \quad (25)$$

Since astronomical observations reveal that the space of the universe is flat, i.e. $k = 0$ in the Robertson-Walker metric ansatz [19], we then get the large scale Lorentz violation modified Friedmann Equation,

$$\left(\frac{\dot{a}}{a}\right)^2 = \frac{\rho + \rho_K}{3} = \frac{\rho + \Lambda_0}{3} - \dot{\mathcal{X}}^2 - 2\dot{\mathcal{X}} \frac{\dot{a}}{a} \quad (26)$$

and

$$\ddot{a} = - \frac{a}{2} \left( p + p_K + \frac{\rho + \rho_K}{3} \right) = - \frac{a}{2} \left( p + \frac{\rho}{3} \right) + \frac{1}{3} a \Lambda_0 - \frac{d}{dt} (a \dot{\mathcal{X}}) \quad (27)$$

From the modified Friedmann equation [27], the condition for accelerating expansion of the universe can be easily obtained as

$$\frac{a}{2} \left( p + \frac{\rho}{3} - 2\Lambda_0 \right) + \frac{d}{dt} (a \dot{\mathcal{X}}) < 0 \quad (27)$$

As discussed in [10], the prediction of Lorentz violation parameters $f_\mu(x)$ needs quantum gravity and the inflation model in some detail rather than the present model. However, it would be suggestive to seek some phenomenological approximations about $\mathcal{X}(t)$ instead of constructing the quantum gravity theory and inflation model in detail at first step. It is inspiring to compare the Friedmann equations of $\Lambda$CDM model,

$$\left(\frac{\dot{a}}{a}\right)^2 = \frac{\rho + \rho_\Lambda}{3} = \frac{\rho + \Lambda}{3} \quad (28)$$

and

$$\ddot{a} = - \frac{a}{2} \left( p + p_\Lambda + \frac{\rho + \rho_\Lambda}{3} \right) = - \frac{a}{2} \left( p + \frac{\rho}{3} \right) + \frac{1}{3} a \Lambda \quad , (29)$$

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with the modified one (26) and (27), where \( \rho_\Lambda = \Lambda, \ p_\Lambda = - \Lambda, \ \Lambda \approx 0.713775 \) and \( H_0 \) is the present Hubble constant. It should be noted that \( \Lambda \) in eq. (28) and (29) is an observation input or a theoretical input by a more fundamental theory. Once the equation of state (EoS), \( p = w \rho, \) of the cosmic media is given, one can make prediction on the evolution of cosmic observables \( a, \ p, \) and \( \rho \) etc.. In the modified case of (26) and (27), one need one more input, \( \mathcal{K} \) which should be given by more fundamental theory, quantum gravity, or also as an observational input phenomenologically just like \( \Lambda \) in eq. (28) and (29) of \( \Lambda \text{CDM} \) model. Since \( \Lambda \text{CDM} \) model works well phenomenologically, we can employ it to fix the evolution of \( \mathcal{K} \) phenomenologically and approximately.

Eq. (26) can be rewritten into

\[
H^2(t) = \frac{\rho + \Lambda_0}{3} - \mathcal{K}^2 - 2 \mathcal{K} H(t)
\]

(30)

the evolution equation for Hubble constant \( H(t) = \frac{\dot{a}(t)}{a(t)} \). So can be eq. (28) as

\[
H^2(t) = \frac{\rho + \Lambda}{3}
\]

(31)

with the solution

\[
H(t) = \sqrt{\frac{\rho + \Lambda}{3}}.
\]

(32)

The solutions of eq. (30) can then be chosen

\[
H(t) = \sqrt{\frac{\rho + \Lambda}{3} - \mathcal{K}(t)}.
\]

(33)

The initial value of \( \mathcal{K}(t) \) can be obtained by the comparison between (26) and (28),

\[
2 \mathcal{K}(t_0) \frac{\dot{a}(t_0)}{a(t_0)} + \mathcal{K}(t_0)^2 = \frac{\Lambda_0}{3} - \frac{\Lambda}{3}
\]

(34)

where \( t_0 \approx H_0^{-1} \) is the moment at present, or the age of universe now and \( H_0 \) is the Hubble constant. Phenomenologically, there are two choices on the initial value of \( \mathcal{K}(t) \),

\[
\mathcal{K}_1 = H_0 \left( \sqrt{1 - \frac{\Lambda - \Lambda_0}{3H_0^2}} - 1 \right)
\]

(35)

and

\[
\mathcal{K}_2 = -H_0 \left( \sqrt{1 - \frac{\Lambda - \Lambda_0}{3H_0^2}} + 1 \right).
\]

(36)

In this scenario, a constrain condition for \( \Lambda_0 \) must be satisfied as

\[
\Lambda_0 \geq -0.401\Lambda \approx -\frac{2}{5}\Lambda
\]

(37)

in order to get a reasonable evolution of \( \mathcal{K}(t) \).

We can make three kind of approximations listed as Case A, B and C to fix the evolution of \( \mathcal{K} \) on the basis of the second Friedmann equation of \( \Lambda \text{CDM} \) model (29).

By comparing eq. (27) and eq. (29), we can make the first approximation named **Case A** as

\[
\frac{d}{dt}(a \mathcal{K}) = -\frac{1}{3} a (\Lambda - \Lambda_0)
\]

(38)

Suppose the EoS of cosmic media is \( p(t) = w(t) \rho(t) \), eq. (26) and eq. (27) can be converted to the dependence of evolution \( a(t) \) on \( \mathcal{K} \) and \( \Lambda_0 \)

\[
\frac{\ddot{a}}{a} + \frac{3w + 1}{2} \frac{\dot{a}^2 + k}{a^2} = -\mathcal{K} - \frac{3w + 1}{2} \mathcal{K}^2 - (3w + 2) \frac{\dot{a}}{a} \mathcal{K} + \frac{w + 1}{2} \Lambda_0
\]

(39)
by eliminating $p$ and $\rho$. Do the same to eq.(28) and eq.(29), we get the dependence of evolution $a(t)$ on $\Lambda$ as

$$\frac{\dot{a}}{a} + \frac{3w + 1}{2} \frac{\ddot{a}^2}{a^2} = \frac{w + 1}{2} \Lambda$$  

We can make the second approximation named **Case B** by requiring

$$\dot{\mathcal{K}} + \frac{3w + 1}{2} \mathcal{K}^2 + (3w + 2) \frac{\dot{a}}{a} \mathcal{K} = \frac{w + 1}{2} (\Lambda_0 - \Lambda),$$

which relates contortion to $w(t)$.

Suppose the dark partner satisfies equation of state $p_k = w_0 \rho_k$, we make the third approximation named **Case C** as

$$(3w_0 + 1) \mathcal{K}^2 + (6w_0 + 4) \frac{\dot{a}}{a} \mathcal{K} + 2 \dot{\mathcal{K}} = (w_0 + 1) \Lambda_0.$$

The evolution of $H(t)$ and $\mathcal{K}(t)$ can be determined by the equations $\mathcal{K}(t) = K_0$ together with one of the approximations eq.(38), eq.(41) and eq.(42) and with the initial conditions $H(t_0) = H_0$ and $\mathcal{K}(t_0) = \mathcal{K}_1$ of eq.(35) or $\mathcal{K}(t_0) = \mathcal{K}_2$ of eq.(36).

Table 1: Models of Approximation in Large Scale Lorentz Violation Cosmology

| Values of $\mathcal{K}(t_0)$ | Evolution Eq. of $\mathcal{K}(t)$ |
|-------------------------------|----------------------------------|
| **Case A1** $\mathcal{K}(t_0) = \mathcal{K}_1$ | $\dot{\mathcal{K}} = \mathcal{K}_1$ |
| **Case A2** $\mathcal{K}(t_0) = \mathcal{K}_2$ | $\dot{\mathcal{K}} = \mathcal{K}_2$ |
| **Case B1** $\mathcal{K}(t_0) = \mathcal{K}_1$ | $\dot{\mathcal{K}} = \mathcal{K}_1$ |
| **Case B2** $\mathcal{K}(t_0) = \mathcal{K}_2$ | $\dot{\mathcal{K}} = \mathcal{K}_2$ |
| **Case C1** $\mathcal{K}(t_0) = \mathcal{K}_1$ | $\dot{\mathcal{K}} = \mathcal{K}_1$ |
| **Case C2** $\mathcal{K}(t_0) = \mathcal{K}_2$ | $\dot{\mathcal{K}} = \mathcal{K}_2$ |

Table 1 summarizes the models of approximation discussed above.

The evolution of $H(t)$ and $\mathcal{K}(t)$ versus $t$ and $\Lambda_0$ in all the cases of approximations are presented in Fig.1 and Fig.2 where we set $w(t) = 0$ in the EoS of the cosmic media for cold matter dominates the energy density in the period of late time expansion of the universe and $w_0 = -0.89$ and $w_0 = -1$ in Case C as examples.

It is apparent that the evolutions of Hubble constant versus $t$ in these models are very close and ones of **Case B** and $\Lambda$CDM model tend to almost coincide with each other in the long time evolution. There is clear but small difference between the longtime evolution curves of **Case A** and $\Lambda$CDM model for $\Lambda_0 < 0$ while one between **Case C** at $w_0 = -1$ and $\Lambda$CDM model is a little bit larger in the initial value $\mathcal{K}(t_0) = H_0 \left( \sqrt{1 - \frac{\Lambda - \Lambda_0}{3H_0^2}} - 1 \right)$ case. The differences among the evolution curve for $H(t)$ decrease with the increase of $\Lambda_0$ so that ones for **Case A, Case B, Case C** and $\Lambda$CDM model are almost the same. In the long time evolution of Hubble constant, the difference between **Case A** and $\Lambda$CDM model decreases as the increase of $\Lambda_0$ while one between **Case C** at $w_0 = -\frac{2}{3}$ and $\Lambda$CDM model is getting larger more and more. The case for $w_0 = -\frac{1}{3}$ of **Case C** is similar to case of $w_0 = -\frac{2}{3}$ while only differes at the deviation from $\Lambda$CDM model is larger than the $w_0 = -\frac{2}{3}$ case.

In the initial value $\mathcal{K}(t_0) = -H_0 \left( \sqrt{1 - \frac{\Lambda - \Lambda_0}{3H_0^2}} + 1 \right)$ case, the difference between the time evolution of **Case A** and $\Lambda$CDM model is always smaller than one between **Case C** and $\Lambda$CDM model no matter what $\Lambda_0$ is. The evolution curves of these several models coincide almost identically where $\Lambda_0$ value is around zero. The deviations of both **Case A** and **Case C** from $\Lambda$CDM model increase along with the increase of $\Lambda_0$ value while one for **Case C** is getting larger than **Case A**.

The evolution of $\mathcal{K}(t)$ versus $t$ and $\Lambda_0$ are presented in Fig.2 where the results of **Case A** and **Case B** are very close but one of **Case C** deviates a little bit larger from one of **Case A** and **Case B**. Especially
Figure 1: The evolution of Hubble constant $H$ versus $t$ in the case of initial value $\mathcal{H}(t_0) = H_0 \left( \pm \sqrt{1 - \frac{\Delta - \Lambda_0}{3H_0^2}} - 1 \right)$ and among three case of approximation and $\Lambda CDM$ model, (a) and (d) $\Lambda_0 = -0.2\Lambda$, (b) and (e) $\Lambda_0 = 0$, (c) and (f) $\Lambda_0 = 0.2\Lambda$
in the initial value $\mathcal{K}(t_0) = H_0 \left( \sqrt{1 - \frac{\Lambda - \Lambda_0}{3H_0^2}} - 1 \right)$ case, results of Case A and Case B are almost the same while one of Case C differs. We can find that Case C is not a good approximation probably for the reason that $w_0$ need to evolve either.

Figure 2: The evolution of $\mathcal{K}$ versus $t$ in the case of initial value $\mathcal{K}(t_0) = H_0 \left( \pm \sqrt{1 - \frac{\Lambda - \Lambda_0}{3H_0^2}} - 1 \right)$ among three case of approximation, (a) and (d) $\Lambda_0 = -0.2\Lambda$, (b) and (e) $\Lambda_0 = 0$, (c) and (f) $\Lambda_0 = 0.2\Lambda$.

In the case of non-zero contortion, the world line of a photon is still the light-like geodesic curve as in the Riemann space-time case rather than the auto-parallel curve. The redshift formula is the same as in the Lorentz invariant zero contortion case,

$$1 + z = \frac{a_0}{a}$$  \hspace{1cm} (43)
and
\[ \frac{dz}{dt} = -a_0/a^2 \]  
(44)

where \( a_0 = a(t_0) \). We can convert \( H \) and \( \mathcal{H} \) versus \( t \) to \( z \),

\[ \dot{H}(t) = H'(z) \left( -\frac{a_0}{a^2} \right) aH = -(1 + z) H(z) H'(z) \]  
(45)

and
\[ \dot{\mathcal{H}}(t) = \frac{d\mathcal{H}}{da} \frac{da}{dt} = \mathcal{H}'(a) aH = -(1 + z) H(z) \mathcal{H}'(z) \]  
(46)

with relations
\[ \dot{H}(t) = \frac{dH}{dt} = \frac{dH}{da} \frac{da}{dt} = H'(a) aH \]  
(47)

and
\[ H'(a) = \frac{dH}{da} = \frac{dH}{dz} \frac{dz}{da} = H'(z) \frac{dz}{da} \]  
(48)

By definition of luminosity distance \( d_L \) (see [18]) we can get
\[ d_L(z) = (1 + z) \int_0^z \frac{1}{H(z')} \, dz' \]  
(49)

in the case of \( k = 0 \) and
\[ \frac{dt}{dz} = -\frac{1}{1 + z} \frac{d}{dz} \left( \frac{d_L}{1 + z} \right) \]  
(50)

With eq. (49) and eq. (50), we can convert Friedmann equations (39) to equation for \( d_L(z) \) and \( \mathcal{H}(z) \) versus redshift \( z \),

\[ \frac{(1 + z)^6 d_L''(z)}{((1 + z) d_L'(z) - d_L(z))^3} - \frac{1}{2} \frac{(1 + z)^4}{((1 + z) d_L'(z) - d_L(z))^2} - \frac{(1 + z)^3 k'(z)}{(1 + z) d_L'(z) - d_L(z)} \]
\[ + \frac{2(1 + z)^2}{(1 + z) d_L'(z) - d_L(z)} \mathcal{H}'(z) = \frac{1}{2} \mathcal{H}^2(z) = \Lambda_0 \]  
(51)

So can we do to three approximations eq. (38), eq. (41) and eq. (42) to get
\[ \frac{(1 + z)^2}{(1 + z) d_L'(z) - d_L(z)} \mathcal{H}'(z) - \frac{(1 + z)^3}{(1 + z) d_L'(z) - d_L(z)} \mathcal{H}'(z) = \frac{1}{3} \left( \Lambda_0 - \Lambda \right) , \]  
(52)

\[ \frac{4(1 + z)^2 \mathcal{H}'(z) - 2(1 + z)^3 K'(z)}{(1 + z) d_L'(z) - d_L(z)} + \mathcal{H}^2(z) = \Lambda_0 - \Lambda \]  
(53)

and
\[ \frac{2(1 + z)^2 ((3w_0 + 2) \mathcal{H}'(z) - (1 + z) \mathcal{H}'(z))}{(1 + z) d_L'(z) - d_L(z)} + (3w_0 + 1) \mathcal{H}^2(z) = (w_0 + 1) \Lambda_0 \]  
(54)

Comparisons of the luminosity distance \( d_L \) curve versus redshift \( z \) among three models of approximation and \( \Lambda \text{CDM} \) model are presented in Fig. 3.

It is observed that no matter what value \( w_0 \) and \( \Lambda_0 \) take, the luminosity distance \( d_L \) versus redshift \( z \) curves of Case B and \( \Lambda \text{CDM} \) model appear to coincide with each other almost. The difference between Case A and \( \Lambda \text{CDM} \) model and one between Case C at \( w_0 = -1 \) and \( \Lambda \text{CDM} \) model decrease with the increase of \( \Lambda_0 \) in the initial value \( \mathcal{H}'(t_0) = H_0 \left( \sqrt{1 - \frac{\Lambda - \Lambda_0}{3H_0^2}} - 1 \right) \) case, while the latter is a little bit larger than the former. The difference between Case C and \( \Lambda \text{CDM} \) model increases with the increase of \( w_0 \) and reaches to a quite large deviation when \( w_0 = -\frac{1}{3} \).

The behavior of luminosity distance \( d_L \) versus redshift \( z \) curves for the initial value \( \mathcal{H}'(t_0) = -H_0 \left( \sqrt{1 - \frac{\Lambda - \Lambda_0}{3H_0^2}} + 1 \right) \) case is similar to one of \( \mathcal{H}'(t_0) = H_0 \left( \sqrt{1 - \frac{\Lambda - \Lambda_0}{3H_0^2}} - 1 \right) \) case.

The distance modulus is defined as \( \mu = 25 + 5 \log_{10}(d_L/Mpc) \).
Figure 3: the comparison of luminosity distance in the case of initial value $\mathcal{K}(t_0) = H_0 \left( \pm \sqrt{1 - \frac{\Lambda_0}{3H_0^2}} - 1 \right)$ among three case of approximation and $\Lambda CDM$ model, (a) and (d) $\Lambda_0 = -0.2\Lambda$, (b) and (e) $\Lambda_0 = 0$, (c) and (f) $\Lambda_0 = 0.2\Lambda$
Figure 4: the comparison between theoretical prediction of the three case of approximation and observation on luminosity distance modulus in the case of initial value $x(t_0) = H_0 \left( \sqrt{1 - \frac{\Lambda - \Lambda_0}{3H_0^2}} - 1 \right)$, in the case of $\Lambda_0 = 0.2\Lambda$. 
4 The Effective Quintessence Potential from String Landscape

By the comparison of eq.(24) and eq.(26), one can introduce the idea of effective cosmological constant which is responsible for the accelerating expansion of the universe rather than the bare one \( \Lambda_0 \),

\[
\Lambda_{\text{eff}}(t) = \Lambda_0 - 3 \left( \mathcal{K}(t)^2 + 2 \mathcal{K}(t) \frac{\dot{a}(t)}{a(t)} \right). 
\]

The bare cosmological constant \( \Lambda_0 \) is just the vacuum energy density which plays only a partial role in the accelerating expansion in our approach. It is the non-trivial contortion which determines the accelerating expansion eventually.

Phenomenologically, the \( \Lambda_{\text{eff}} \) can be regarded as an energy density produced by some auxiliary fields which are responsible for the accelerating expansion such as quintessence field etc\[20\]. We can consider the action for gravity in the form of

\[
S_q = \int d^4x \sqrt{-g} \left[ \frac{1}{2} M_{\text{pl}}^2 R - \frac{1}{2} g^{\mu\nu} \partial_\mu \phi \partial_\nu \phi - V(\phi) \right].
\]

The corresponding energy density for the field \( \phi \) is

\[
\rho_\phi = \frac{\dot{\phi}^2}{2} + V(\phi) 
\]

and the pressure of field \( \phi \) is

\[
p_\phi = \frac{\dot{\phi}^2}{2} - V(\phi) 
\]

The continuity equation for the \( \phi \) field is

\[
\ddot{\phi} + 3H \dot{\phi} + V_{,\phi} = 0 
\]

The solution of eq.\(59\) would give the evolution \( \phi(t) \). In the analog of LSLV model by quintessence, we have

\[
\Lambda_{\text{eff}}(t) = \frac{\dot{\phi}^2(t)}{2} + V(\phi(t)) 
\]

Then we have the relation

\[
V(\phi(t)) = \Lambda_{\text{eff}} + \frac{\dot{\Lambda}_{\text{eff}}}{6H} 
\]

The evolution of \( V(\phi(t)) \) given by \( \Lambda_{\text{eff}}(t) \) versus \( t \) bifurcates at some critical value of \( \Lambda_0 \), named as \( \Lambda_{\text{crit}} \), i.e. \( V(\phi(t)) \) decreases monotonically along with the increase of \( t \) when \( \Lambda_0 \leq \Lambda_{\text{crit}} \) while its evolution has a local minimum when \( \Lambda_0 > \Lambda_{\text{crit}} \). Table\[2\] summarizes the value of \( \Lambda_{\text{crit}} \) for all the cases of approximation in consideration.

We find that the Case C in the \( \mathcal{K}' \)’s first initial value case, \( \mathcal{K}(t_0) = H_0 \left( \sqrt{1 - \frac{\Lambda - \Lambda_0}{3H_0^2}} - 1 \right) \), \( V(\phi(t)) \) keeps decreasing monotonically along with the increase of \( t \) for all value of \( \Lambda_0 \) allowed when \( w_0 > -\frac{8}{9} \), i.e. the solution of \( \Lambda_{\text{crit}} \) does not exit in this case. However, the comparison of theoretical prediction and observation on luminosity distance modulus reveals Case C1 when \( w_0 > -\frac{8}{9} \) can be ruled out by observation on luminosity distance modulus versus redshift curve as is shown in Fig\[4\]. It is reasonable to argue that \( w_0 > -\frac{8}{9} \) for Case C1 is not a good approximation to fix the evolution of \( \mathcal{K} \).

It can be observed that the solution of \( \Lambda_{\text{crit}} \) values are all around \( \Lambda_0 = 0 \), the division of string landscape and swampland. We can make a reasonable analysis about the deviation of the \( \Lambda_{\text{crit}} \) values from \( \Lambda_0 = 0 \). All the \( \Lambda_{\text{crit}} \) values in Table\[2\] are obtained with the approximation on the evolution of \( \mathcal{K} \), we can guess \( \Lambda_{\text{crit}} \) would be exactly zero in a more elaborated model of the evolution of \( \mathcal{K} \). The monotonic \( V(\phi) \) on \( \phi \) will result a monotonic \( V(\phi(t)) \) versus \( t \). On the other hand, the appearance of local minimum in \( V(\phi(t)) \) corresponds to a non-monotonic \( V(\phi) \) on \( \phi \). It reveals that \( \Lambda_{\text{crit}} \) may be the real division of string landscape and swampland.
Figure 5: the $\Lambda_{\text{crit}}$ solutions in all the cases with both the initial values $\mathcal{H}(t_0) = H_0 \left( \pm \sqrt{1 - \frac{\Lambda - \Lambda_0}{3H_0^2}} - 1 \right)$
Table 2: The critical values of $\Lambda_0$ triggers the transition from monotonic evolution of $V(\phi(t))$ to the one with local minimum in different models

| Case          | Initial Values of $X(t)$                           | Critical Values of $\Lambda_0$ |
|---------------|----------------------------------------------------|--------------------------------|
| Case A        | $X(t_0) = H_0 \left( \frac{1 - \frac{X - X_0}{3H_0^2}}{1 - \frac{X - X_0}{3H_0^2}} - 1 \right)$ | -0.05$A$                      |
|               | $X(t_0) = -H_0 \left( \frac{1 - \frac{X - X_0}{3H_0^2}}{1 - \frac{X - X_0}{3H_0^2}} + 1 \right)$ | -0.187$A$                     |
| Case B        | $X(t_0) = H_0 \left( \frac{1 - \frac{X - X_0}{3H_0^2}}{1 - \frac{X - X_0}{3H_0^2}} - 1 \right)$ | -0.071$A$                      |
|               | $X(t_0) = -H_0 \left( \frac{1 - \frac{X - X_0}{3H_0^2}}{1 - \frac{X - X_0}{3H_0^2}} + 1 \right)$ | -0.2144$A$                    |
| Case C($w_0 = -1$) | $X(t_0) = H_0 \left( \frac{1 - \frac{X - X_0}{3H_0^2}}{1 - \frac{X - X_0}{3H_0^2}} - 1 \right)$ | 0.00001$A$                    |
|               | $X(t_0) = -H_0 \left( \frac{1 - \frac{X - X_0}{3H_0^2}}{1 - \frac{X - X_0}{3H_0^2}} + 1 \right)$ | 0.00001$A$                    |
| Case C($w_0 = -\frac{8}{9}$) | $X(t_0) = H_0 \left( \frac{1 - \frac{X - X_0}{3H_0^2}}{1 - \frac{X - X_0}{3H_0^2}} - 1 \right)$ | 0.119$A$                      |
|               | $X(t_0) = -H_0 \left( \frac{1 - \frac{X - X_0}{3H_0^2}}{1 - \frac{X - X_0}{3H_0^2}} + 1 \right)$ | 0.075$A$                      |
| Case C($w_0 = -\frac{7}{9}$) | $X(t_0) = H_0 \left( \frac{1 - \frac{X - X_0}{3H_0^2}}{1 - \frac{X - X_0}{3H_0^2}} - 1 \right)$ | none                          |
|               | $X(t_0) = -H_0 \left( \frac{1 - \frac{X - X_0}{3H_0^2}}{1 - \frac{X - X_0}{3H_0^2}} + 1 \right)$ | 0.143$A$                      |
| Case C($w_0 = -\frac{2}{3}$) | $X(t_0) = H_0 \left( \frac{1 - \frac{X - X_0}{3H_0^2}}{1 - \frac{X - X_0}{3H_0^2}} - 1 \right)$ | none                          |
|               | $X(t_0) = -H_0 \left( \frac{1 - \frac{X - X_0}{3H_0^2}}{1 - \frac{X - X_0}{3H_0^2}} + 1 \right)$ | 0.2$A$                       |
| Case C($w_0 = -\frac{1}{3}$) | $X(t_0) = H_0 \left( \frac{1 - \frac{X - X_0}{3H_0^2}}{1 - \frac{X - X_0}{3H_0^2}} - 1 \right)$ | none                          |
|               | $X(t_0) = -H_0 \left( \frac{1 - \frac{X - X_0}{3H_0^2}}{1 - \frac{X - X_0}{3H_0^2}} + 1 \right)$ | 0.306$A$                      |
To account for the inflation and the accelerating expansion, the AdS vacua of string landscape need to be lifted to dS ones by some unnatural mechanisms such as the KKLT or the LVS construction\cite{2, 3} etc. However, the two criteria on dS swampland conjecture rule out the meta-stable dS types vacua and leave only the possibility of quintessence like vacua. The result obtained in this paper reveals that large scale Lorentz violation can lift the AdS vacua to quintessence like ones effectively.

5 Summary and Outlook

The KKLT or the LVS construction on uplifting the AdS vacua to dS ones relies on carefully adding $\bar{D}_3$-branes into the compactification. The $\bar{D}_3$-branes tension in a sufficiently warped background, in the presence of quantum corrections, can be a small enough correction to lift the formerly AdS vacuum to positive cosmological constant, without destabilizing the minimum. However, there are also many cases with no-go theorems to uplift the AdS vacua to dS ones. Several results reveal no-go theorems on dS vacua in string theory constructions, with restrictions on ingredients used in string theory, typically specific combinations of fluxes, D-branes, orientifolds, etc\cite{4}. Moreover, the second criterion on swampland conjecture exclude the effective theory with meta-stable dS vacua as theories with UV completion. It seems that to account for the accelerating expansion and inflation with a simple $\Lambda CDM$ model with positive cosmological constant is difficult. If the positive cosmological constant comes from vacuum energy density of a vacuum, both the construction of such vacuum is fragile and it will belong to the swampland. However, the quintessence like potential can pass the restriction by the second criterion of the swampland conjecture, which will also contribute a decreasing positive vacuum energy density.

Our result obtained in this paper shows another route to uplift the AdS vacua to an effective quintessence types vacuum energy which can satisfy the second criterion of the swampland conjecture and without adding extra ingredients such as fluxes or D-branes into the commodification. The result only relies on the consideration of large scale Lorentz violation frozen by the inflation from the quantum gravity. Actually, the origin causing accelerating expansion is assorted to the effect of quantum gravity in this scenario. The construction has a UV completion in this sense.

Our approach is also different from quintessence model which has a scalar mode as a part of the gravity at large scale. The quintessence scalar field is either elementary or an effective theory of other elementary physics mechanisms, which need either detection examination or clarification on the relation from other mechanism. In our approach, the quintessence field is actually an effective description of the dark partner contribution from contortion and it is not necessary but inspiring.

We employ three kinds of approximation to fix the evolution and magnitude of $\mathcal{K}(t)$. However, a more fundamental approach should be given later. We would start with a specific model of quantum gravity and a inflation model to achieve a model with a stronger ability of prediction on the evolution and magnitude of $\mathcal{K}(t)$. The Lorentz violation from quantum gravity can be traced along with the inflation and the frozen large scale Lorentz violation can be predicted hence. The present approach relies on gauge principle with the local symmetry group as the proper subgroup of Lorentz group. Actually, the Lorentz violation can be other types rather than the totally breakdown of some symmetry generators. The Lorentz violation can be achieved with a boost transformation differing from Lorentz boost only. Some discussions assert that the quantum gravity effect can modify the Poincaré algebra into a deformed one, a Hopf algebra e.g. $\kappa$-Poincaré etc\cite{21}. The gauge principle is not applicable in these cases. How to introduce the Lorentz violation in long range effective gravity beyond gauge principle is under investigation.

In the most approaches on the extension of general relativity with torsion, the torsion or the contortion can not propagate and can only be nontrivial in the region of matter source distribution with spin\cite{22}. The large scale spacetime must be torsion free and so is the case for the universe. The space-time felt by matter motion described by the left-hand side of eq.\cite{13} is the Riemann-Cartan space-time which is determined not only by energy-momentum but also by spin as in eq.\cite{4}. In our approach, though the torsion caused by frozen large scale Lorentz violation propagates neither, it distributes as the shadow of matter distribution and evolves along as the evolution of the universe. Matter moves in a Riemannian spacetime instead of Riemann-Cartan one and interacts with the distribution of torsion in the usual way of gravitation, i.e. the effective energy momentum tensor contributed by contortion distribution participates the determination of spacetime curvature and matter moves along the geodesic line and feels
the contortion effect through the spacetime curvature. However, the spin of matter particle can interact with the effective spin angular momentum tensor directly in addition to the gravitational interaction. The large scale long propagation of particle with spin may exhibit the deviation from geodesic line and causes the advance or delay arrival time variation with its energy. There are indeed some indication of such events for the gamma-ray bursts and neutrinos and the delay differs between gamma-ray and neutrino\[23\]. A detailed investigation on the subject based on our approach of frozen large scale Lorentz violation is performing.

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References

[1] C. Vafa, *The String landscape and the swampland*, hep-th/0509212.

[2] S. Kachru, R. Kallosh, A. D. Linde and S. P. Trivedi, *De Sitter vacua in string theory*, Phys. Rev. D68 (2003) 046005 hep-th/0301240.

[3] M. Cicoli, J. P. Conlon and F. Quevedo, *General Analysis of LARGE Volume Scenarios with String Loop Moduli Stabilisation*, JHEP 10 (2008) 105 0805.1029.

[4] T. D. Brennan, F. Carta and C. Vafa, *The String Landscape, the Swampland, and the Missing Corner*, PoS TASI2017 (2017) 015 1711.00864.

[5] G. Obied, H. Ooguri, L. Spodyneiko and C. Vafa, *De Sitter Space and the Swampland*, 1806.08362.

[6] P. Agrawal, G. Obied, P. J. Steinhardt and C. Vafa, *On the Cosmological Implications of the String Swampland*, Phys. Lett. B784 (2018) 271 1806.09718.

[7] H. Ooguri, E. Palti, G. Shiu and C. Vafa, *Distance and de Sitter Conjectures on the Swampland*, Phys. Lett. B788 (2019) 180 1810.05506.

[8] S. M. Carroll, J. A. Harvey and othes, *Noncommutative field theory and lorentz violation*, Physical Review Letters 87 (2001) 141601 hep-th/00105082.

[9] J. Alfaro, H. A. Morales-Técotl, M. Reyes and L. Urrutia, *Alternative approaches to lorentz violation invariance in loop quantum gravity inspired models*, Physical Review D 70 (2004) 084002 gr-qc/0404113.

[10] J. Shen and X. Xue, *Large Scale Lorentz Violation Gravity and Dark Energy*, 1802.03502.

[11] L. Perivolaropoulos, *Large scale cosmological anomalies and inhomogeneous dark energy*, Galaxies 2 (2014) 22 1401.5044.

[12] J. Martin, *Everything you always wanted to know about the cosmological constant problem (but were afraid to ask)*, Comptes Rendus Physique 13 (2012) 566 1205.3365.

[13] Y. Wu, X. Xue, L. Yang and T.-C. Yuan, *The effective gravitational theory at large scale with lorentz violation*, 1510.00514.

[14] Y. Wu and X. Xue, *SIM(2) gravity gauge theory*, Journal of East China Normal University (Natural Science) 2016 (2016) 76.

[15] L. Yang, Y. Wu, W. Wei, X. Xue and T.-C. Yuan, *The effective gravitation theory at large scale with lorentz violation*, Chinese Science Bulletin (2017) 944.
[16] W. Wei, J. Shen, Y. Wu, L. Yang, X. Xue and T.-C. Yuan, \textit{E(2) gauge theory model of effective gravitational theory at large scale}, \textit{Acta. Phys. Sin.} 66 (2017) 130301.

[17] R. Aldrovandi and J. G. Pereira, \textit{Teleparallel Gravity}, vol. 173. Springer, Dordrecht, 2013, 10.1007/978-94-007-5143-9.

[18] S. Weinberg, \textit{Cosmology}. 2008.

[19] J. T. Nielsen, A. Guffanti and S. Sarkar, \textit{Marginal evidence for cosmic acceleration from type ia supernovae}, \textit{Scientific reports} 6 (2016) 35596 1506.01354.

[20] S. Tsujikawa, \textit{Quintessence: A Review}, \textit{Class. Quant. Grav.} 30 (2013) 214003 1304.1961.

[21] L. Smolin, \textit{Classical paradoxes of locality and their possible quantum resolutions in deformed special relativity}, \textit{Gen. Rel. Grav.} 43 (2011) 3671 1004.0664.

[22] F. W. Hehl, P. Von Der Heyde, G. D. Kerlick and J. M. Nester, \textit{General Relativity with Spin and Torsion: Foundations and Prospects}, \textit{Rev. Mod. Phys.} 48 (1976) 393.

[23] X. Zhang and B.-Q. Ma, \textit{Testing Lorentz invariance and CPT symmetry using gamma-ray burst neutrinos}, \textit{Phys. Rev.} D99 (2019) 043013 1810.03571.