THE SLOAN DIGITAL SKY SURVEY/XMM-NEWTON QUASAR SURVEY: CORRELATION BETWEEN X-RAY SPECTRAL SLOPE AND EDDINGTON RATIO

G. Risaliti1,2, M. Young1,3, and M. Elvis1
1 Harvard-Smithsonian Center for Astrophysics, 60 Garden Street, Cambridge, MA 02138, USA
2 INAF-Osservatorio Astrofisico di Arcetri, Largo E. Fermi 5, I-50125 Firenze, Italy
3 Boston University, Astronomy Department, 725 Commonwealth Avenue, Boston, MA 02215, USA

ABSTRACT

We present a correlation between the 2–10 keV spectral slope $\Gamma_X$ and the Eddington ratio $L/L_{\text{EDD}}$ in a sample of $\sim$400 Sloan Digital Sky Survey (SDSS) quasars with available hard X-ray spectra from XMM-Newton serendipitous observations. We find that the $\Gamma_X - L/L_{\text{EDD}}$ correlation is strongest in objects with black hole (BH) masses determined from the H$\beta$ line, and weaker (but still present) for those based on Mg II. An empirical nonlinear correction of the Mg II-based masses, obtained by comparing the mass estimates in SDSS quasars having both H$\beta$ and Mg II measurements, significantly increases the strength of the correlation. No correlation is found among objects with BH masses derived from C IV, confirming that this line is not a reliable indicator of the BH mass. No significant correlation is found with the bolometric luminosity, while a $\Gamma_X - M_{\text{BH}}$ relation is present, though with a lower statistical significance than between $\Gamma_X$ and $L/L_{\text{EDD}}$. Our results imply a physical link between the accretion efficiency in the (cold) accretion disk of active galactic nuclei and the physical status of the (hot) corona responsible for the X-ray emission.

Key words: galaxies: active – X-rays: galaxies

1. INTRODUCTION

The origin of the X-ray emission in quasars is not well understood. In the widely accepted disk–corona model (Haardt & Maraschi 1993), the X-rays are produced in a hot phase (the corona) reprocessing the primary optical/UVM emission of the disk. However, the mechanisms of energy transfer to the hot phase, and its geometry and size are not clear. Therefore, it is not known how the basic physical parameters (such as black hole (BH) mass, accretion rate, total luminosity) affect the X-ray emission. Laor et al. (1997) found a correlation between the soft X-ray (0.2–2 keV) slope and the FWHM of the H$\beta$ emission line in a sample of 23 low-redshift quasars, and suggested that the physical parameter driving the correlation is the Eddington ratio, $L/L_{\text{EDD}}$, where $L$ is the bolometric luminosity. Further studies in the 2–10 keV energy band (e.g., Brandt et al. 1997; Shemmer et al. 2006) confirmed this correlation for the hard X-ray slope, and is therefore not related to the active galactic nucleus (AGN) “soft excess.” In particular, a $\Gamma_X - L/L_{\text{EDD}}$ correlation has been found in a sample of $\sim$150 Sloan Digital Sky Survey (SDSS) quasars with Chandra spectra, and spanning a large redshift range (Kelly et al. 2008). Recently, Shemmer et al. (2008) presented a similar analysis based on a sample of 35 quasars, spanning more than 3 orders of magnitude in luminosity, and obtained a stronger correlation between $\Gamma_X$ and the Eddington ratio $L/L_{\text{EDD}}$ than with FWHM(H$\beta$), breaking the partial degeneracy between these two quantities. The main limitation of past studies involving X-ray slopes is that either they rely on low-quality X-ray data, or they are based on relatively small samples. The availability of large, homogeneous samples of quasars with high-quality optical and X-ray spectral data is now opening new possibilities in this field.

In particular, we recently obtained a new sample of $\sim$800 quasars within the SDSS/XMM-Newton survey (Risaliti & Elvis 2005; Young et al. 2009, hereafter Y09) obtained by cross-correlating the SDSS DR5 quasar catalog with the XMM-Newton public archive. Since we only selected serendipitous XMM-Newton observations, this sample can be considered as randomly extracted from the SDSS quasars. About 500 quasars in this sample have good enough X-ray data to perform a basic spectral analysis and obtain a continuum slope. A complete description of the survey, with an analysis of the general properties of the sample, is presented in Y09.

The simultaneous availability of optical/UV and X-ray spectra not only allows a more precise analysis of the $\alpha_{\text{OX}}$–luminosity correlation (this subject is developed in a companion paper M. Young et al. 2009, in preparation), but also opens a whole new field of analysis of the possible correlations among X-ray spectral parameters, optical/UV continuum and line emission, and physical parameters of the AGN.

Here we extend the analysis of the correlation between $\Gamma_X$ and optically derived and global quantities such as $M_{\text{BH}}$, $L/L_{\text{EDD}}$, and $L$ for our sample of SDSS–XMM-Newton quasars, consisting of more than 400 objects.

2. THE SAMPLE

Our sample is taken from the SDSS/XMM-Newton quasar survey (Y09). Since we are interested in a homogeneous analysis of X-ray slopes, we excluded radio-loud (RL) quasars and broad absorption line (BAL) quasars, whose X-ray spectra may be contaminated by synchrotron emission (in RLs) or affected by heavy absorption (in BALs). Since a complete removal of BALs is possible only for sources at $z > 1.5$, and $z < 1.5$, there may be some residual BAL contamination in our sample. By removing objects with flat X-ray spectra (see below), we expect to further clean the sample from BALs, which are typically harder in X-rays than non-BAL quasars (e.g., Gallagher et al. 2006). The remaining spurious objects are expected to be $\lesssim 5\%$ of the sample.

For the 403 quasars with a signal-to-noise ratio ($S/N$) $> 6$, we used the available X-ray and optical/UV spectral data to obtain estimates of the relevant physical parameters for our analysis: the X-ray continuum slope, $\Gamma_X$, the BH mass, $M_{\text{BH}}$, and the bolometric luminosity, $L_{\text{BOL}}$. The sample spans about 3 orders
of magnitude in optical luminosity (this can be estimated from the bottom panel of Figure 1, where we show the bolometric luminosities, which are roughly proportional to the optical ones, as explained below). The 2–10 keV luminosities range from $10^{45}$ to $10^{45.5}$ erg s$^{-1}$ for most sources, with small tails down to $10^{42}$ erg s$^{-1}$ at $z < 0.5$, and up to $10^{46}$ erg s$^{-1}$ at $z > 1.5$. More details on the luminosity distribution can be found in Y09.

The X-ray slopes $\Gamma_X$ were obtained from basic power law plus Galactic absorption models applied to the spectral data in the rest-frame 2–10 keV band. The data reduction was performed with the SAS package following the standard recommended steps in the XMM-Newton Science Data Center Web site. The analysis was made with the Sherpa analysis package. All the details about the data reduction and analysis are presented in Y09. The sample spans an X-ray luminosity $10^{43}$ erg s$^{-1} < L_X < 5 \times 10^{46}$ erg s$^{-1}$, and a redshift range $0.1 < z < 4.5$.

The BH masses have been estimated (Shen et al. 2008, hereafter S08) from the widths of the optical broad emission lines and the underlying continuum luminosity. Three different optical lines are used, depending on the redshift of the sources: H$\beta$ ($0 < z < 0.9$), Mg $\Pi$ (0.4 < $z < 2.2$), and C IV (1.7 < $z < 4.5$). When the BH masses from two different lines are available, we prefer H$\beta$ over Mg $\Pi$, and Mg $\Pi$ over C IV (see Section 4 for more details on this issue). For the H$\beta$ and Mg $\Pi$ lines, the widths were obtained through a two-Gaussian fit, after subtracting a template reproducing the Fe II emission. The adopted templates are those of Boroson & Green (1992) for the H$\beta$ line and of Salviander et al. (2007) for the Mg $\Pi$ line. For the C IV line, the continuum has been fitted with a simple power law, while the line FWHM has been estimated from the analytic three-Gaussian fit described in Laor et al. (1994). Full details on the optical spectral analysis are described in S08.

The bolometric luminosities are derived from the continuum monochromatic luminosity at the line wavelength, adopting a correction based on an analytical intrinsic spectral energy distribution (SED). The optical to UV SED is approximated by three power laws with spectral indices (in the $v - f_v$ plane) $\alpha_1 = -2$ in the 1–10 $\mu$m interval, $\alpha_2 = -0.44$ from 1$\mu$m to 1200 Å, and $\alpha_3 = -1.76$ from 1200 Å to 500 Å, in agreement with the average SED of SDSS quasars (Richards et al. 2006). The X-ray luminosity is modeled with a power-law continuum with the observed photon index, starting at 0.1 keV, and an exponential decrease with a cutoff energy $E_C = 100$ keV. The X-ray to optical ratio is directly obtained from the observed data.

We analyzed the correlation between the X-ray photon index, $\Gamma_X$, and (1) the Eddington ratio, $L/L_{\text{edd}}$, (2) the bolometric luminosity, $L_{\text{bol}}$, and (3) the BH mass, $M_{\text{BH}}$. The main results for the whole sample are shown in Figure 1. As mentioned above, our data consist of three different subsamples, depending on the line used to estimate the BH mass. There are at least two reasons to perform a separate analysis for each subsample: (1) the BH mass–line width correlation is directly calibrated using reverberation mapping on H$\beta$ widths (e.g., Peterson et al. 2004), while it is indirectly calibrated using H$\beta$ for the other lines (Vestergaard & Peterson 2006); (2) the luminosity and redshift ranges spanned by the three lines are quite different (Figure 1, bottom panel), and a possible luminosity dependence has to be considered. We therefore repeated the same analysis for the total sample and each of the three subsamples.

### 3. Statistical Analysis

Our statistical analysis consists of several steps and checks. We first performed a nonparametric Spearman rank test on each correlation, followed by a least-squares linear fit. In order to take into account the effect of possible deviant points, we performed

---

4. http://xmm.esac.esa.int/sas/8.0.0/
5. http://cxc.harvard.edu/sherpa/
6. This part gives a negligible contribution, and represents the Rayleigh–Jeans tail of the accretion disk emission. The actual observed emission at wavelengths $\lambda > 1 \mu$m is due to reprocessing by dust, and therefore in not included here.

---

**Figure 1.** Correlations between $\Gamma_X$ and $L/L_{\text{edd}}$, $M_{\text{BH}}$, and $L_{\text{bol}}$, for our sample. The error bars are not shown for clarity. A typical error is shown in the top panel. The three colors refer to the line used to estimate the BH mass. The lines show the best linear fits and the dispersions. The statistical errors on the slopes are small with respect to the dispersions, as shown in Table 1, and are not plotted here.
the above analysis with different selections based on the quality of X-ray data.

1. We tried different cuts in the S/N of the X-ray spectra ($6 < S/N < 10$), and in the quality of the fit of the X-ray data ($0.5 < \chi^2 < 2$). No significant dependence on these choices has been found. The final analysis was performed on objects with $S/N > 8$ in the X-ray spectra (see Y09 for details) and $\chi^2 < 1.5$ in the X-ray fit.

2. We excluded possibly “bad” points, i.e., those with extreme best-fit values of $\Gamma_X$ ($\Gamma_X > 3$ or $\Gamma_X < 1$), or large deviations from the best fit in a statistical sense ($|\langle \Delta \Gamma_X \rangle |/\Gamma_X > 5$). Again, the results remain consistent for all the different cuts.

3. We estimated the errors for the slope $m$ and intercept $q$ through a bootstrap analysis, consisting of repeating the linear fits on randomly selected sets of data drawn from our sample, allowing for repetitions. This technique takes into account strongly deviating points that could affect the correlations.

All these checks demonstrate that our results are stable and do not depend on single deviating points.

The results for the total sample are shown in Table 1. The final sample contains 343 objects. The subsamples consist of 82 (H$\beta$), 290 (Mg $\Pi$), and 58 (CIV) data points. 58 objects have both H$\beta$ and Mg $\Pi$ measurements, while 29 have both Mg $\Pi$ and CIV.

4. RESULTS

We found a highly statistically significant correlation (probability of null correlation $P < 0.1\%$) between the X-ray slope and the Eddington ratio $L/L_{\text{EDD}}$, and a weaker, but still significant anticorrelation between $\Gamma_X$ and the BH mass. No significant trend ($P > 5\%$) is found with the bolometric luminosity (Table 1).

Separate analysis of the subsamples show that the $\Gamma_X-L/L_{\text{EDD}}$ correlation is strongest for objects whose $M_{\text{BH}}$ is determined from H$\beta$ (Figure 2), weaker for Mg $\Pi$ objects, and absent for CIV objects. The linear correlation coefficient is higher than 0.5, conventionally considered the threshold for a “strong” correlation, only for the $\Gamma_X-L/L_{\text{EDD}}$ correlation in the H$\beta$ group.

Since $L/L_{\text{EDD}} \propto M_{\text{BH}}^{-1}$, and $M_{\text{BH}} \propto \text{FWHM}^2 L^\alpha$ ($\alpha = 0.52 \pm 0.04$ in the calibration of Bentz et al. 2006), a strong $\Gamma_X-L/L_{\text{EDD}}$ correlation may be due to an intrinsic correlation between $\Gamma_X$ and FWHM. We tested this possibility and found a highly significant $\Gamma_X-\text{FWHM}$ correlation (slightly weaker than the $\Gamma_X-L/L_{\text{EDD}}$ one in the total and H$\beta$ samples). As a more general approach, we tested correlations of the form
FWHM$^{-2} \times L^{\beta}$ varying the exponent $\beta$ in the interval 0–0.8, and we always found the same level of correlation for any value of $\beta$ (null correlation probability $P < 10^{-8}$, linear correlation coefficient $r$ in the range 0.53–0.60). This confirms that the luminosity term does not contribute to the observed correlation. This may be due to either a physical independence of these quantities, or to a too narrow luminosity range in our sample (Figure 1). We note that Shemmer et al. (2008) claim a stronger dependence of $\Gamma_X$ on $L/L_{Edd}$ than on FWHM(H$\beta$) for their quasar sample, which is smaller (35 objects) but spanning a larger luminosity interval.

The $\Gamma_X$–log($M_{BH}$) relation is statistically significant (Table 1), but rather weak, with a slope of $\sim$0.3. Since $L_{Edd} \propto M_{BH}$, some degree of correlation of $\Gamma_X$ with $M_{BH}$ is expected, given the strength of the $\Gamma_X$–$L/L_{Edd}$ correlation. Formally, it is not possible to remove this partial degeneracy. We only note that the stronger $\Gamma_X$–$L/L_{Edd}$ correlation suggests that this is the physically relevant relation, and, therefore, that there is no independent support for a direct physical dependence of $\Gamma_X$ on $M_{BH}$.

The decreasing strength (in terms of slope of a linear correlation) and statistical significance of the correlation going from H$\beta$– Mg $\uparrow$ to C IV–estimated FWHM BH masses could be due to the different average luminosities of the three subsamples (Figure 1, bottom panel), or could be due to a problem related to the estimates of $M_{BH}$. We directly explored both possibilities.

1. **Luminosity effects.** The luminosity intervals of the three subsamples overlap significantly. We analyzed the $\Gamma_X$–$L/L_{Edd}$ correlation for luminosity-limited subsamples of the Mg $\uparrow$ group, matched to the luminosity distribution of the H$\beta$ group. We performed Pearson’s tests and linear correlation fits for Mg $\uparrow$ objects with bolometric luminosities log($L_{BOL}$) < 45.5 (23 objects), log($L_{BOL}$) < 45.8 (76 objects), and log($L_{BOL}$) < 46 (123 objects). In no case did we find a better correlation than in the whole Mg $\uparrow$ sample. This suggests that luminosity is not the main reason for the observed trend.

2. **BH mass estimates.** We compared the BH mass estimates from pairs of lines in the S08 sample. As discussed in S08, about 8000 objects in their sample have BH mass estimates from both H$\beta$ and Mg $\uparrow$, and about 5000 from both Mg $\uparrow$ and C IV. In Figure 3, we plot the mass estimates for the two groups (this plot is analogous to Figure 6 in S08), with a best-fit linear correlation. We find a strong (though not linear) correlation between H$\beta$ and Mg $\uparrow$ masses, while no significant correlation is found between Mg $\uparrow$ and C IV masses. These two findings have two different possible explanations.

(a) Both H$\beta$ and Mg $\uparrow$ are expected to be produced by virialized gas, and are therefore in principle good mass indicators. The mismatch between H$\beta$ and Mg $\uparrow$ masses is likely related to the uncertainties in the measurement of the Mg $\uparrow$ line parameters: the line is a doublet, and the estimate of the strength of the contaminating Fe $\uparrow$ emission is more problematic than for the H$\beta$ line. Recently, Onken & Kollmeier (2008) discussed this issue and found that a correction has to be applied to the average Mg $\uparrow$-based mass distribution, in order to make it compatible with the one based on H$\beta$. In our case, in order to apply a correction to the Mg $\uparrow$-based values of $M_{BH}$ in single sources, we used the best-fit linear correlation shown in Figure 3 (upper panel) (log($M_{BH}$(H$\beta$)) = 1.8 × log($M_{BH}$(Mg $\uparrow$)) – 6.8), under the assumption that the H$\beta$-based masses are correct. This should improve the precision of the estimates of $M_{BH}$, at least in a statistical sense. We repeated the same statistical analysis on the so-modified Mg $\uparrow$ subsample, and we obtained a stronger correlation (probability of null correlation < 10$^{-6}$), and a linear correlation coefficient $r = 0.3$. This is still lower than that found for H$\beta$ sample, as expected, since our correction is only statistical, and a significant residual dispersion is still present between the H$\beta$-based and Mg $\uparrow$-based estimates (Figure 3, upper panel). However, the strength of the correlation is significantly higher than that found with the original $M_{BH}$ values ($r = 0.21$, Table 1).

(b) C IV is a high ionization line, produced in the inner broad line region by gas probably having nonvirialized components (for example, associated with an outflow). This makes this line a poor estimator of the BH mass (Baskin & Laor 2005; Netzer et al. 2007). The lack of correlation between Mg $\uparrow$ and C IV masses confirms this finding, and provides an explanation for the absence of correlations between the X-ray

![Figure 3](image-url)
spectral slope and the Eddington ratio and the BH mass.

5. CONCLUSIONS

We presented a strong correlation between the X-ray photon index, $\Gamma_X$, and the Eddington ratio, $L/L_{\text{EDD}}$ for a sample of $\sim 400$ AGNs having good quality X-ray and optical spectra from the SDSS quasar survey and serendipitous $\text{XMM-Newton}$ observations. A weaker, but statistically significant, correlation between $\Gamma_X$ and the BH mass $M_{\text{BH}}$ was also found.

The correlations found here are important in two respects: first, a $\Gamma_X-L/L_{\text{EDD}}$ correlation is an important constraint for theoretical emission models. In general, it suggests a strong link between the accretion rate and the physical conditions in the hot corona producing the X-rays. For example, recently Cao (2009) showed that a $\Gamma_X-L/L_{\text{EDD}}$ relation, together with the $\alpha_{\text{OX}}$–luminosity relation, depend on the balance between the magnetic field and the gas and electron pressures in a disk–corona model where the corona is heated by magnetic field reconnections.

Second, as already pointed out by Shemmer et al. (2008), our results show that the X-ray slope can be used as an Eddington ratio estimator. Since the X-ray luminosity can be converted to a total luminosity through bolometric corrections and the $\alpha_{\text{OX}}$–luminosity correlation, an estimate of the BH mass can be obtained from the X-ray data alone. Given the high dispersion of the correlation (0.4 dex), this is not at present a reliable method for single sources. However, it could be a powerful technique when used to estimate average values in large samples, such as the ones which are expected to be available in the future from new missions like E-Rosita.

We are currently working at expanding the SDSS–$\text{XMM-Newton}$ quasar survey, using the DR7 version of the SDSS$^7$ and the 2009 $\text{XMM-Newton}$ archive. This new sample will have about twice as many quasars with good quality X-ray spectra, and will allow further investigations of the multi-wavelength correlations among AGNs.

We are grateful to the referee for his/her constructive comments. This work has been partly supported by grants prin-miur 2006025203, ASI-INAF I/088/06/0, and NASA NNX07AI22G.

REFERENCES

Baskin, A., & Laor, A. 2005, MNRAS, 356, 1029
Bentz, M. C., Peterson, B. M., Pogge, R. W., Vestergaard, M., & Onken, C. A. 2006, ApJ, 644, 133
Boroson, T. A., & Green, R. F. 1992, ApJS, 80, 109
Brandt, W. N., Mathur, S., & Elvis, M. 1997, MNRAS, 285, L25
Cao, X. 2009, MNRAS, 394, 207
Gallagher, S. C., Brandt, W. N., Chartas, G., Pródenes, R., Garmire, G. P., & Sánchez-Ramírez, R. M. 2006, ApJ, 644, 709
Haardt, F., & Maraschi, L. 1993, ApJ, 413, 507
Kelly, B. C., Bechtold, J., Trump, J. R., Vestergaard, M., & Siemiginowska, A. 2008, ApJS, 176, 355
Laor, A., Bahcall, J. N., Jannuzi, B. T., Schneider, D. P., Green, R. F., & Hartig, G. F. 1994, ApJ, 420, 110
Laor, A., Fiore, F., Elvis, M., Wilkes, B. J., & McDowell, J. C. 1997, ApJ, 477, 93
Netzer, H., Lira, P., Trakhtenbrot, B., Shemmer, O., & Cury, I. 2007, ApJ, 671, 1256
Onken, C. A., & Kollmeier, J. A. 2008, ApJ, 689, L13
Peterson, B. M., et al. 2004, ApJ, 613, 682
Richards, G. T., et al. 2006, ApJS, 166, 470
Risaliti, G., & Elvis, M. 2005, ApJ, 629, L17
Salviander, S., Shields, G. A., Gebhardt, K., & Bonning, E. W. 2007, ApJ, 662, 131
Shemmer, O., Brandt, W. N., Netzer, H., Maiolino, R., & Kaspi, S. 2006, ApJ, 646, L29
Shemmer, O., Brandt, W. N., Netzer, H., Maiolino, R., & Kaspi, S. 2008, ApJ, 682, 81
Shen, Y., Greene, J. E., Strauss, M. A., Richards, G. T., & Schneider, D. P. 2008, ApJ, 680, 169 (S08)
Vestergaard, M., & Peterson, B. M. 2006, ApJ, 641, 689
Young, M., Elvis, M., & Risaliti, G. 2009, ApJS, 183, 17

$^7$ http://www.sdss.org/DR7/