LETTER TO THE EDITOR

Tracking Advanced Planetary Systems (TAPAS) with HARPS-N. *

IV. TYC 3667-1280-1 - the most massive red giant star hosting a warm Jupiter.

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ABSTRACT

Context. We present the latest result of the TAPAS project that is devoted to intense monitoring of planetary candidates that are identified within the PennState-Toruń planet search.

Aims. We aim to detect planetary systems around evolved stars to be able to build sound statistics on the frequency and intrinsic nature of these systems, and to deliver in-depth studies of selected planetary systems with evidence of star-planet interaction processes.

Methods. The paper is based on precise radial velocity measurements: 13 epochs collected over 1920 days with the Hobby-Eberly Telescope and its High-Resolution Spectrograph, and 22 epochs of ultra-precise HARPS-N data collected over 961 days.

Results. We present a warm-Jupiter (\(T_{\text{eq}} = 1350\text{ K}\), \(m_\text{e} = 5.4 \pm 0.4 \text{ M}_\oplus\) companion with an orbital period of 26.468 days in a circular \((e = 0.036)\) orbit around a giant evolved \((\log g = 3.11 \pm 0.09, R = 6.26 \pm 0.86 R_\odot)\) star with \(M_\star = 1.87 \pm 0.17 M_\odot\). This is the most massive and oldest star found to be hosting a close-in giant planet. Its proximity to its host \((a = 0.21 \text{ au})\) means that the planet has a 13.9 \pm 2.0% probability of transits; this calls for photometric follow-up study.

Conclusions. This massive warm Jupiter with a near circular orbit around an evolved massive star can help set constraints on general migration mechanisms for warm Jupiters and, given its high equilibrium temperature, can help test energy deposition models in hot Jupiters.

Key words. Stars: evolution, activity, late-type, planetary systems; Planets and satellites: detection, individual: TYC 3667-1280-1 b; Planet-star interactions.

1. Introduction

The population of hot Jupiters (HJ), Jupiter-mass planets on short-period orbits \((P < 10 \text{ days}, \text{ or within } 0.1 \text{ au})\), which were unveiled with the finding of 51 Peg b (Mayor & Queloz 1995), was probably one of the least expected discoveries in search for exoplanets. The origin of this rare class of objects, only known to be present around approximately 0.5–1% of FGK stars (Wright et al. 2012; Howard et al. 2012), generally involves early (e.g., Type II), or late migration scenarios (following planet-planet scattering or secular perturbations from more distant objects) since formation beyond the snow line is required. Recently, however, the in-situ formation of these systems has also been proposed (Boley et al. 2016).

Warm Jupiters (WJ, Fogg & Nelson 2009), Jupiter-mass planets on 10–100 day orbits, are even more intriguing. WJ have observed eccentricities too low for significant tidal evolution, and while the missing HJ from the evolved star-planet population can be naturally explained by engulfment by the star, the lack so far of detected WJ requires further explanation (see, e.g.,
Villaver et al. [2014]. Recently, Frewen & Hansen [2016] invoked Kozai-Lidov oscillations to predict an entire removal of WJ planets by the time the star reaches $R > 5 \ M_\odot$, while an identical constant eccentricity population survives beyond $4 \ M_\odot$. The authors argued that the WJ migrate through Kozai-Lidov oscillations. Dawson & Murray-Clay [2013] presented evidence that both smooth disk migration and dynamical interactions may play a role in WJ formation, with the efficiency depending on the host’s metallicity. Dong et al. [2014] showed that high-eccentricity WJ are more likely to have an external Jovian perturber and migrate through dynamical interactions. This finding was recently supported by Huang et al. [2016], who found that different from HJ, half of the WJ have companions. The mutual inclination of orbits in six such WJ-external perturber pairs reported by Dawson & Chang [2014] indeed points to Kozai-Lidov oscillations as a mechanism responsible for the inward migration of WJ.

We here present the discovery of a warm-Jupiter planet ($P = 26,468.0 \pm 0.7 \text{d}$) orbiting an evolved star with $R = 6.3 \ M_\odot$ in a circular orbit. Clearly, this unique system can help set constraints on the invoked mechanism for WJ formation through the Kozai-Lidov oscillations. Furthermore, its proximity to a very luminous star means that it has the equivalence temperature of a hot Jupiter. If hot-Jupiter inflation works by depositing irradiation into the planet’s deep interiors, then planetary radii should increase in response to the increased irradiation of this evolved star, containing structure theories on the physics of the hot-Jupiter inflation mechanism (Lopez & Fortney 2013). Moreover, there is a high probability of transit detection of this unique object, which may then serve as a perfect laboratory for planet formation and migration theories.

The paper is organized as follows: in Sect. 3 we present the available observations and discuss the influence of the stellar activity on the radial velocity (RV) variation measurements. Section 4 shows the results of the Keplerian data modeling, in Sect. 5 we discuss the transit possibilities, and in Sect. 6 we discuss the results and present the conclusions.

2. Observations, radial velocities, and activity

TYC 3667-1280-1 (2MASS J00513296+5825342) belongs to a sample of about 300 planetary or brown dwarf (BD) candidates identified in the complete sample of over 1000 stars for RV variations with the 9.2 m Hobby-Eberly Telescope (HET; Ramsey et al. 1996) and its High-Resolution Spectrograph (HRS, Tall 1998) since 2004 within the Penn State - Toruń Centre for Astronomy Planet Search (PTPS; Niedzielski et al. 2007, 2016a) selected for a more intense precise RV follow-up within Tracking Advanced Planetary Systems (TAPAS) with HARP-S-N (Niedzielski et al. 2015, Adamow et al. 2015, Niedzielski et al. 2016). A summary of the available data for TYC 3667-1280-1 is given in Table 1.

The spectroscopic observations presented here were made with the HET and its HRS in the queue-scheduled mode (Shetrone et al. 2007) and with the 3.58 m Telescopio Nazionale Galileo (TNG) and its High-Accuracy Radial velocity Planet Searcher in the Northern hemisphere (HARPS-N, Cosentino 2012).

For HET HRS spectra we used a combined gas-cell (Marcy & Butler 1992, Butler et al. 1996), and cross-correlation (Queloz 1995, Pepe et al. 2002) method for precise RV and spectral line bisector (BS) measurements, respectively. The implementation of this technique to our data is described in Nowak et al. (2012) and Nowak et al. (2013).

| Parameter | Value | Parameter | Value |
|-----------|-------|-----------|-------|
| V (mag)   | 9.86 ± 0.02 (1) | $M/M_{\odot}$ | 1.87 ± 0.17 (7) |
| B-V (mag) | 1.00 ± 0.06 (1) | log $L/L_{\odot}$ | 1.38 ± 0.10 (7) |
| (B-V)$_0$ (mag) | 0.84 (2) | $R/R_{\odot}$ | 6.26 ± 0.86 (7) |
| $M_\star$ (mag) | 0.99 (2) | log age (yr) | 9.13 ± 0.10 (7) |
| d (pc) | 481 ± 37 (6) | $v_{\text{rad}}$ (m s$^{-1}$) | 3.00 ± 0.1(7) |
| $T_\text{eff}$ (K) | 5130 ± 24 (2) | $P_{\text{rad}}$ (d) | 0.075 ± 0.002 (7) |
| log g | 3.11 ± 0.09 (2) | $P_{\text{inst}}$ (d) | 99 ± 19 (7) |
| [Fe/H] | $-0.08 ± 0.05 (2)$ | $v_{\text{rad}}$ $\sin i_\star$ (km s$^{-1}$) | 3.2 ± 0.4 (7) |
| RV (km s$^{-1}$) | 9.54 ± 0.06 (3) | $\Delta L(\text{NLTE})$ | 0.80 ± 0.21 (4) |

Notes. References: (1) [Hog et al. 2000], (2) [Zielinski et al. 2012], (3) [Nowaki 2012], (4) [Adamow et al. 2014], (5) [Adamczyk et al. 2016], (6) calculated from $M_\star$, (7) This work.

3. Keplerian analysis

Keplerian orbital parameters were derived using a hybrid approach (e.g., Goździewski et al. 2003, Goździewski & Migaszewski 2006, Goździewski et al. 2007), in which the PIKAIA-based, global genetic algorithm (GA; Charbonneau 1995) was combined with MPFit algorithm (Markwardt 2009) to find the best-fit Keplerian orbit delivered by RVLIN [Wright & Howard 2009] modified to allow the stellar jitter to be fitted as a free parameter (Ford & Gregory 2007, Johnson et al. 2011). The RV bootstrapping method (Murdoch et al. 1993, Kuerster et al. 1997, Marcy et al. 2005, Wright et al. 2007) was employed.
where $T_\text{f}$ is the mid-transit time for an epoch $E$ (in HJD).

Detecting a transit event would allow determining the planetary radius, orbital inclination, planetary mass, and mean density. The planet would be a very interesting object for studies of planet-star interactions in late stages of stellar evolution. A non-detection would set constraints on the orbital inclination or planetary radius.

### 5. Discussion and conclusions

Given the orbital period of the companion (26.468 days) and the stellar mass ($M = 1.87 \pm 0.17 M_\odot$), stellar radius ($R = 6.26 \pm 0.86 R_\odot$), and the evolutionary stage of the host ($\log g = 3.11 \pm 0.09$), the planetary system TYC 3667-1280-1 is certainly unique. Its host is an evolved A star entering the red giant branch. The planet resides at only $7.2 \pm 1.0 R_\star$ and is on the brink of engulfment, which will occur when its host reaches the radius of $\sim 20 R_\star$ in $\sim 1.7 \times 10^8$ yr. This fate is only marginally delayed using a different mass-loss prescription (see Villaver et al. 2014).

It is the most compact planetary system ($a = 0.21$ au) hosted by an evolved ($\log g \leq 3.5$), intermediate-mass ($M_\star \geq 1.5 M_\odot$) star. The only planetary systems discovered using RV method that bear a resemblance to TYC 3667-1280-1 are HIP 67851 b (Jones et al. 2015), HD 102956 b (Johnson et al. 2010), and 8 U Mi b (Lee et al. 2015). However, while the host star of HIP 67851 b has a similar mass of 1.63 $\pm$ 0.22 $M_\odot$ and is at the similar stage of stellar evolution ($\log g = 3.2 \pm 0.2$), the planet itself resides on a much wider orbit with $a = 0.459 \pm 0.021$ au and is much lighter (1.38 $\pm$ 0.15 $M_\oplus$). HD 102956 b is even less massive (0.96 $\pm$ 0.05 $M_\oplus$) and is hosted by a less evolved ($\log g = 3.5 \pm 0.06$), 1.68 $\pm$ 0.11 $M_\odot$ star, similarly to 8 U Mi b, a 1.5 $\pm$ 0.2 $M_\odot$ planet orbiting a 1.8 $\pm$ 0.1 $M_\odot$, log $g = 2.57 \pm 0.03$ star in 0.49 $\pm$ 0.03 au orbit (Fig. 2).

TYC 3667-1280-1 b is the most massive planet of these, which makes it a strong candidate for the study of tidal dissipation, similar to WASP-18b (Hellier et al. 2009). The proximity of its evolved luminous host TYC 3667-1280-1 b causes its equilibrium temperature to be very high and places the planet among the hottest Jupiters known (see Fig. 3). This makes it a perfect laboratory for studying inflation processes (Lopez & Fortney 2016).

The relatively high probability of transits adds to its attractiveness. The estimated probability of $13.9 \pm 2.0 \%$ is high enough to make this worthwhile, even though the expected transit depth is relatively low (0.3–1.1 mmag) and the transit duration may reach up to 29 hours. TYC 3667-1280-1 b certainly deserves more attention and a photometric follow-up in search for transits.

Furthermore, the planetary system of TYC 3667-1280-1 represents an excellent test case on which to study the formation of WJ through Kozai-Lidov oscillations because its mere existence, along with its almost circular orbit, defies the theories of migration through this mechanism. Planets like TYC 3667-1280-1 b are extremely rare. One object in the complete sample of $\sim 1000$ PTPS stars suggests a frequency of only $\sim 0.1%$. However, given one object detected so far within the total PTPS sample of 103 stars with $M_\star \geq 1.5 M_\odot$, we can estimate that objects like this appear around $\sim 1%$ of evolved intermediate-mass stars in our sample. They are apparently as common as HJ around the main-sequence stars. Our discovery suggests that WJ around intermediate-mass stars survive the evolution of their hosts at the beginning of the red giant branch, where TYC 3667-1280-1 is currently located.
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### Table 3. HET and HRS RV and BS measurements (m s$^{-1}$) of TYC 3667-1280-1

| MJD     | RV   | $\sigma_{RV}$ | BS  | $\sigma_{BS}$ |
|---------|------|---------------|-----|---------------|
| 54394.155075 | 258.5 | 10.3 | -48.4 | 14.4 |
| 54398.137124 | 217.8 | 9.3 | -36.9 | 24.0 |
| 54748.181551 | -45.9 | 10.5 | -20.8 | 31.5 |
| 54779.129988 | -185.2 | 10.0 | -40.3 | 26.9 |
| 55184.154144 | 156.9 | 9.8 | -44.6 | 24.1 |
| 55469.395590 | -202.0 | 8.9 | -83.8 | 19.7 |
| 55500.319306 | 12.1 | 10.6 | 17.5 | 34.0 |
| 55895.253142 | -67.2 | 8.8 | -28.3 | 21.9 |
| 56127.428623 | -122.8 | 10.6 | -79.8 | 22.8 |
| 56232.306667 | -68.9 | 8.7 | -0.4 | 19.4 |
| 56280.170359 | 241.9 | 10.2 | -90.6 | 27.1 |
| 56295.121470 | 80.6 | 8.7 | -10.7 | 27.9 |
| 56314.096829 | -188.1 | 9.2 | -66.5 | 24.8 |

### Table 4. TNG and HARPS-N RV and BS measurements (m s$^{-1}$) of TYC 3667-1280-1

| MJD     | RV   | $\sigma_{RV}$ | BS  |
|---------|------|---------------|-----|
| 56277.0056838 | 5852.76 | 0.52 | -208.5 |
| 56293.9512124 | 5497.29 | 0.89 | -83.2 |
| 56320.904931 | 5522.61 | 1.44 | -63.3 |
| 56560.1855872 | 5609.82 | 1.86 | -150.9 |
| 56647.0042254 | 5788.52 | 2.42 | -70.0 |
| 56684.8799384 | 5372.04 | 2.27 | 30.2 |
| 56895.1655851 | 5412.09 | 2.56 | -98.3 |
| 56926.9657616 | 5440.58 | 2.93 | 19.5 |
| 56927.2172308 | 5453.96 | 2.99 | 16.2 |
| 56969.8405909 | 5640.65 | 1.84 | -137.5 |
| 56970.050176 | 5635.57 | 3.56 | -187.8 |
| 56970.070991 | 5609.47 | 4.10 | -168.5 |
| 56991.9991155 | 5817.72 | 2.75 | -154.6 |
| 57034.819799 | 5545.53 | 2.58 | 53.8 |
| 57034.9211894 | 5553.18 | 2.77 | 57.9 |
| 57034.9663153 | 5555.87 | 2.78 | 65.6 |
| 57065.8310694 | 5803.94 | 1.99 | 16.8 |
| 57065.8540671 | 5804.62 | 3.24 | 14.9 |
| 57196.209713 | 5661.05 | 1.55 | -95.9 |
| 57196.2327686 | 5726.66 | 2.09 | 141.2 |
| 57238.1151227 | 5426.42 | 1.63 | -103.7 |
| 57238.2091972 | 5418.56 | 1.79 | -60.3 |