C/O AND O/H RATIOS SUGGEST SOME HOT JUPITERS ORIGINATE BEYOND THE SNOW LINE

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ABSTRACT

The elemental compositions of planet-hosting stars serve as proxies for the primordial compositions of protoplanetary disks within which the planets form. The temperature profile of the disk governs the condensation fronts of various compounds, and although these chemically distinct regions migrate and mix during the disk lifetime, they can still leave an imprint on the compositions of the forming planets. Observable atmospheric compositions of hot Jupiters, when compared against their host stars, could potentially constrain their formation and migration processes. We compared the measured planetary and stellar abundances of carbon and oxygen for 10 systems with hot Jupiters. If the planets formed by core accretion with significant planetesimal accretion and migrated through the disk, the hot Jupiter atmospheres should be substantially super-stellar in O/H and substellar in C/O. On the contrary, however, we find that currently reported abundances of hot Jupiters have generally super-stellar C/O ratios, although present uncertainties on the reported O/H and C/O ratios are too large to reach a firm conclusion. Improved measurements from the James Webb Space Telescope will enable more precise measurements for more hot Jupiters, and we predict, based on the current marginal trend, that a sizable fraction of hot Jupiters will show enrichment of C/O over and lower O/H than their hosts, similar to HD 209458b.

Key words: planets and satellites: atmospheres – planets and satellites: composition – planets and satellites: formation – stars: abundances

1. INTRODUCTION

Before the detection of 51 Peg b (Mayor & Queloz 1995), there was little evidence to support migration in planet formation models. In the core accretion model described by Pollack et al. (1996), the giant planets form outside the snow line and begin to accrete significant gas envelopes once the cores reach about 10 M_J. Terrestrial planet formation in this model can occur closer to the star because these planets do not accrete appreciable gas envelopes and so are not constrained by gas-disk lifetimes. The discovery of a Jupiter-mass planet on a 4 day orbit provided data to support theories of planet formation models, but the details of core formation change such that they no longer need to form outside the ice line. One notable consequence of this change is where the planet accretes its final gas envelope.

As the number of planetary systems has increased, so have the questions about planet formation models. The difficulty in growing dust beyond centimeter-sized pebbles and the presence of tightly packed systems of rocky planets occupying the same orbital regions as hot Jupiter systems have prompted a re-examination of some underlying assumptions. Recently, there have been two new theories of in situ formation for hot Jupiters (Batygin et al. 2015; Boley et al. 2016). These are still core accretion models, but the details of core formation change such that they no longer need to form outside the ice line. One notable consequence of this change is where the planet accretes its final gas envelope.

During the evolution of the protoplanetary disk, the sequential condensation of solids results in changing abundance ratios for the remaining gas (Öberg et al. 2011; Moriarty et al. 2014). Thus, the location in the disk where a giant planet accretes its gaseous envelope can leave an imprint on its atmospheric composition. Although the evolution of the disk combined with uncertainty in migration mechanism hinder locating a planet’s exact birthplace (Madhusudhan et al. 2014a), certain regions of the disk are chemically distinct. Inside the water snow line, small amounts of oxygen will be bound up in magnesium silicates, but otherwise the gas disk should have a nearly stellar oxygen composition. Between the water ice line and the CO2 ice line at ∼10 au, oxygen is bound up in water ice while carbon condensation is not energetically favored. Thus, the gas in the region is oxygen-poor but carbon-rich, and the C/O ratio is significantly enhanced over its primordial and stellar abundance ratio (Öberg et al. 2011; Piso et al. 2015). Between the CO2 and CO snow line at about 30 au, twice as much oxygen as carbon condenses out, further decreasing the O/H ratio and enhancing the C/O ratio of the gas. Beyond the CO snow line, carbon and oxygen will be in solids reflecting the stellar abundance patterns and the gas will be depleted of any metals.

The location in the disk where the planet accretes the majority of its mass and the fraction of solid versus gas accretion will therefore influence the atmospheric composition of gas-giant planets. This natal imprint then gives us a clear discriminant in the case of hot Jupiter formation and subsequent disk-free migration from beyond the snow line. If the C/O ratio of the hot Jupiter atmosphere is significantly elevated over that of its host star, then it was born beyond the water snow line. Gas-giant planets with elevated C/O ratios and [O/H] less than the host star likely accreted their atmospheres beyond the water ice line and then experienced disk-free migration. In contrast, gas-giant planets with [O/H] comparable to or greater than their host star must have accumulated a significant fraction of solids or migrated early through the disk (Madhusudhan et al. 2014a).

In the atmospheres of Jupiter and Saturn, the carbon abundance is enhanced by a factor of 3.5–7 over that in the
solar photosphere (Pontoppidan et al. 2014, p. 363). However, due to their cold temperatures, measuring their water (i.e., oxygen) abundance requires direct sampling (Wong et al. 2004). The situation is quite different for highly irradiated hot Jupiters where water has been clearly detected in 10 planets (Kreidberg et al. 2014, 2015; Line et al. 2014; Madhusudhan et al. 2014b).

We compare the new stellar [O/H] and C/O abundances of Brewer et al. (2016) to literature C/O and H$_2$O/H$_2$ values of hot Jupiter atmospheres and discuss current limitations and future prospects.

2. CARBON AND OXYGEN

The small number of unblended atomic lines of oxygen and carbon in optical spectra has made it challenging to derive accurate stellar abundances for C and O. For planets, the task is further complicated by the high contrast ratios between planet and star, non-equilibrium chemistry, and clouds. In addition the hot Jupiter must transit in order to even obtain planetary atmospheric spectra. Recent progress in measuring C and O for both stars and hot Jupiters is allowing us to move forward in our understanding of planet formation.

2.1. Hot Jupiter Atmosphere Data

Secondary eclipse spectroscopy from a combination of ground- and space-based observatories gives us information about the dayside atmospheres of hot Jupiters. The homogeneously analyzed catalog of Line et al. (2014) provides C/O ratios and H$_2$O/H$_2$ for nine exoplanet atmospheres and includes 1σ confidence regions. An earlier study by Madhusudhan (2012) collected the results for six planet atmospheres, using a χ$^2$ metric to distinguish between a carbon-rich and an oxygen-rich model for each planet. One of the planets, HD 149026b, is smaller than the others at roughly a Saturn mass although the arguments for its atmospheric composition are the same and so we include it and refer to it as a hot Jupiter for the sake of completeness.

In addition to the eclipse measurements, transmission spectroscopy is available for several planets (Kreidberg et al. 2014, 2015; Madhusudhan et al. 2014b). These measurements give us information about the atmosphere at the day–night terminator of the planet. The planets studied overlap with those observed with eclipse spectroscopy and two of the planet atmospheres were analyzed using both eclipse and transit data to derive self-consistent atmospheric models (Kreidberg et al. 2014, 2015).

Seven of the planets have atmospheric H$_2$O/H$_2$ measurements. Since N(H) ≫ N(O), the hydrogen in the water will be negligible and we can estimate the O/H ratio of the planets as $\sim$0.5 × H$_2$O/H$_2$. A comparison of the planet O/H ratio to that of its host star provides a constraint on how the O/C ratio was enhanced or depleted. This in turn could potentially inform us about the formation conditions of the planetary atmosphere (Öberg et al. 2011).

2.2. Host Star Data

The stellar properties catalog of Brewer et al. (2016) contains abundances for 15 elements including carbon and oxygen. The authors note that corrected trends exist over some subsets of the temperature ranges and that these regions should only be used with caution in comparative analyses. For [C/H] and [O/H], the data are relatively free of trends below 6100 K. This affects both WASP-12 at 6154 K and WASP-14 at 6459 K, although we have chosen to keep WASP-12 in this study as it lies just at the edge of the acceptable range. Nine stars remain in common between the stellar host measurements and the hosts of those with planet atmosphere measurements: HD 189733, HD 209458, HD149026, WASP-12, WASP-19, TrES-2, TrES-3, XO-1, and CoRoT-2. Brewer et al. (2016) derived empirical errors for each of their elements, with [C/H] having an error of 0.026 dex and [O/H] an error of 0.036 dex, resulting in a typical error in C/O of $\sim$10% and [O/H] of $\sim$8%.

3. RESULTS

3.1. Implications of Stellar Abundances for Hot Jupiter Compositions

The C/O and [O/H] ratios of the planet-hosting stars in our sample are shown in Figure 1, based on the stellar abundance data of Brewer et al. (2016) and the subsample of cool dwarfs with well measured carbon and oxygen in Brewer & Fischer (2016). We find two important trends in the distribution of elemental abundances of planet-hosting stars, which have important consequences for exoplanetary compositions. First, the distribution of C/O ratios across our sample shows that the Sun is moderately carbon-rich. Most stars, including planet-hosting stars, have lower C/O ratios (i.e., are more oxygen-rich) compared to the Sun, i.e., C/O < 0.54. In particular, the 10 stars hosting hot Jupiters considered in this work have subsolar C/O ratios. Second, we also find that most of the planet-hosting stars, including most of the hosts of the hot Jupiters in this study, also have super-solar O/H ratios. This follows the previously known general trend that the occurrence rate of giant exoplanets increases with host stellar metallicity, [Fe/H] (Fischer & Valenti 2005), and hence with the O/H ratio, since we find that [O/H] scales with [Fe/H] (Figure 2). These trends suggest that the formation environments of giant exoplanets might be predominantly oxygen-rich.

3.2. Oxygen-rich Core Accretion

The subsolar C/O ratios and super-solar oxygen abundance of most planet-hosting stars lead to observable consequences for hot Jupiter compositions. Under the standard model of giant planet formation with core accretion, the metallicity of the planetary atmosphere is driven largely by the composition of the solids accreted in planetesimals during formation and migration through the disk. Giant planets forming in such an environment could then be expected to be even more enriched in oxygen, i.e., [O/H] ≫ 0 and C/O < 0.54. The resultant atmosphere is expected to be enhanced in all the condensible elements relative to the host star (Pollack et al. 1996; Owen et al. 1999; Atreya et al. 2009; Mouis et al. 2009, 2012), as has been observed for C, N, S, P, and inert gases in Jupiter’s atmosphere (Owen et al. 1999; Atreya et al. 2016). Since oxygen is the largest constituent of planetesimals, primarily via H$_2$O ice and silicates, the atmosphere is particularly enriched in oxygen. Here it is assumed that the elemental composition of the nebula and hence the protoplanetary disk mimics the stellar composition.

The oxygen-rich elemental compositions thus predicted for hot Jupiters imply that H$_2$O is the most abundant volatile in their atmospheres, and is also the most observable molecule with current instruments (Line et al. 2014; Madhusudhan...
et al. 2016). It is well known that in H₂-dominated atmospheres with O-rich compositions (e.g., C/O $\approx$ 0.8) H₂O is the most dominant O-bearing molecule in the observable atmosphere at all temperatures ($\sim$300–3500 K). Thus for hot Jupiters with super-solar O/H and subsolar C/O, the H₂O abundance should be super-solar, with volume mixing ratios H₂O/H₂ $\gtrsim$ 10⁻³ (Madhusudhan 2012); at high temperatures ($\gtrsim$1200 K) a fraction of the O (≤50%) can also be bound in CO. Such abundances of H₂O are easily observable in transmission and emission spectra of hot Jupiters.

If, however, hot Jupiters are found to be underabundant in H₂O, that would imply different formation and migration conditions than those assumed in the standard picture discussed above. In general, it would indicate a paucity of planetesimals accreted during the formation and/or disk-free migration, e.g., via dynamical encounters (Madhusudhan et al. 2014a). In what follows, we compare constraints on atmospheric abundances of several hot Jupiters to those of their host stars to constrain the different formation and migration pathways.

### 3.3. Planetary Versus Stellar Abundances

In this section we investigate how the C/O ratios of hot Jupiters compare with the C/O ratios of their host stars. Typically, constraints on atmospheric properties of transiting hot Jupiters are obtained using transmission spectra of the day–night terminator, when the planet transits in front of the host star, and/or thermal emission spectra from the dayside atmosphere, when the planet is at secondary eclipse. Given an observed spectrum, the atmospheric chemical compositions and 1D averaged temperature profile are derived using atmospheric retrieval methods (see, e.g., reviews by

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**Figure 1.** C/O and [O/H] ratios of the 849 cool dwarfs from Brewer & Fischer (2016) and the hosts with measured planet atmosphere abundances used in this work. The gray dashed lines mark the solar values of C/O and [O/H]. Stars without detected planets are designated by pale blue points, known planet hosts are in orange, and known hot Jupiter hosts are in red. The known hot Jupiter hosts that also have measured C/O and [O/H] of the planet atmosphere are represented by a black star along with their names.

**Figure 2.** Stellar oxygen vs. iron abundance for the 849 cool dwarfs from Brewer & Fischer (2016) and the hosts with measured planet atmosphere abundances used in this work. Colors are the same as in Figure 1. In general, the oxygen abundance is tightly correlated with the metallicity as measured by [Fe/H]. Since hot Jupiters are preferentially found around metal-rich stars (Fischer & Valenti 2005), their hosts will also be oxygen-rich.
Madhusudhan et al. 2014c, p. 739; 2016). In planetary atmospheres, the dominant elements are all in molecular form, due to which the elemental abundances are derived from the respective major molecular species. In particular, the C/O ratios are derived from abundances of H2O, CO, CH4, and, to some extent, CO2, which are the most prominent O- and C-bearing species in H2-rich atmospheres under a wide range of atmospheric conditions (Madhusudhan 2012; Moses et al. 2013; Madhusudhan et al. 2016).

Reliable constraints on C/O ratios of hot Jupiters are possible only when multiband atmospheric observations are available at wavelengths that contain strong spectral features of H2O, CO, and CH4. Currently, such data, and hence constraints on the C/O ratios, are available for several hot Jupiters primarily from thermal emission photometry and/or spectra of their dayside atmospheres obtained using the Spitzer and Hubble space telescopes. In the present study we use the C/O ratios of such systems reported in recent literature (Madhusudhan 2012; Line et al. 2014). On the other hand, several recent studies have also reported precise atmospheric observations of transmission spectra but such data only contain constraints on the H2O abundances. In principle, the H2O abundances can be used to derive constraints on the C/O ratios under certain assumptions of chemical and radiative equilibrium (see, e.g., Benneke 2015), but we do not use those C/O ratios in the present study in the interest of being free from model assumptions.

We compared stellar C/O ratios obtained using the stellar abundance data of Brewer et al. (2016) with the hot Jupiter C/O ratios from the atmospheric retrieval methods of Line et al. (2014) and Madhusudhan (2012). Figure 3 shows the planetary versus stellar C/O ratios for nine systems, although only two have best-fit measurements for planetary C/O with no published uncertainties. The most precisely measured exoplanet atmosphere is that of HD 209458b, which has a C/O ratio 2.2 times that of its host star. All but two of the remaining also have super-stellar C/O ratios. The large uncertainties in the planetary atmosphere measurements mean that all planets except HD 209458b are consistent with having the same C/O ratios as their hosts. Using only the measurements with reported uncertainties in planet C/O, we calculated the mean planet-to-star ratio to be 1.9 ±0.14, indicating that the planetary C/O ratios tend to be higher than their host stars. This suggests that while the stars themselves are oxygen-rich, i.e., have low C/O ratios as discussed in Section 3.1, the planets are preferentially oxygen-poor or carbon-rich.

We also investigate trends in the O/H ratios of hot Jupiters versus their parent stars. To this end we use H2O measurements reported for seven hot Jupiters, using emission and/or transmission spectra, from which we derive the O/H ratio assuming H2O is the most dominant oxygen carrier (Section 3.2). Figure 4 shows the comparisons of the stellar versus planetary O/H ratios. Four of the seven planets have O/H ratios consistent with their host stars. The Line et al. (2014) measurement of HD 189733b shows a marginally super-stellar O/H ratio, although the Madhusudhan (2012) measurement from the transmission spectra is almost three orders of magnitude smaller but still consistent with its host. Both HD 149026b and WASP-19b have lower oxygen abundances than their host stars, although the measurements also have very large uncertainties. HD 209458b shows an unambiguous detection of substellar O/H with both transmission and emission spectroscopy. Combined with its super-stellar C/O ratio, this planet is consistent with formation beyond the snow line and subsequent disk-free migration (Madhusudhan et al. 2014a).

Our results above are critically reliant on the accuracy of the derived molecular abundances in hot Jupiter atmospheres. Although our preliminary analysis above suggests that hot Jupiters are generally oxygen-poor relative to their host stars, more detailed observations are required to improve their abundance constraints. Additionally, recent studies have
suggested the possibility of clouds/hazes in some hot Jupiters which could in principle influence the abundance determinations, particularly from transmission spectra where clouds/hazes are expected to have the most impact (Sing et al. 2015; Madhusudhan et al. 2016). For example, in some cases the low H2O abundances derived from transmission spectra in the near-infrared using the HST WFC3 instrument could potentially be due to clouds/hazes partially obscuring the atmospheres (Madhusudhan et al. 2014b). However, this is less likely to be the case in thermal emission spectra which probe the very hot dayside atmosphere of the planet and for which the observations are in the mid-infrared where clouds/hazes are expected to have the least impact. Overall, however, the fact that our conclusions are consistent between both the transmission and emission abundance estimates suggest that hot Jupiters might indeed be oxygen-poor relative to their host stars.

4. DISCUSSION

Elemental abundance ratios of the host stars provide a proxy for the primordial abundances in the protoplanetary disk and, hence, for the initial chemical conditions of planetary formation. Several recent studies have investigated the range of outcomes of giant planetary compositions depending on their formation locations and migration paths (Oberg et al. 2011; Helling et al. 2014; Madhusudhan et al. 2014a; Piso et al. 2015; Mordasini et al. 2016). It is generally assumed that the atmospheric composition reflects the accretion history of the planet (Pinhas et al. 2016).

The key parameters determining the atmospheric composition of a hot Jupiter are the location and amounts of solids accreted relative to the gas in forming the giant planet. For a solar-composition disk, with a C/O ratio of 0.5, oxygen-rich species dominate the solid composition (Madhusudhan et al. 2011; Mousis et al. 2011, 2012; Johnson et al. 2012; Marboeuf et al. 2014; Thiabaud et al. 2015). Within the H2O snow line (≤2 au), the gas is expected to be of nearly stellar composition since the only condensates are silicates which contain ~20% of the O. Farther out in the disk, the gas and solid compositions are governed by the locations relative to the major snow lines; the gas becomes progressively C-rich, reaching a C/O ratio of ~1 beyond the CO2 snow line, and the solids become progressively O-rich, reaching a stellar C/O ratio beyond the CO snow line, where nearly all volatiles are condensed out. Therefore, hot Jupiters forming within the H2O snow line are expected to be of nearly stellar composition. On the other hand, hot Jupiters forming beyond the H2O snow line with significant planetesimal accretion are always expected to be O-rich, with a maximum C/O equal to that of the nebula when formed beyond the CO snow line.

The planet-host stars in our sample are even more oxygen-rich, with higher [O/H] and lower C/O ratios, compared to the Sun. In the standard core accretion scenario, the corresponding hot Jupiters can be expected to be substantially super-solar in oxygen, and hence in H2O abundances. Such oxygen enrichment in hot Jupiters is readily detectable with current and future instruments. Therefore, if hot Jupiters are found to have substellar oxygen, it would indicate low planetesimal accretion, formation beyond the H2O snow line, and a migration path that inhibits solid accretion, e.g., by dynamical scattering mechanisms rather than migration through the disk (Madhusudhan et al. 2014a).

The precision of most hot Jupiter atmosphere C/O measurements is not currently high enough to rule out in situ formation models or migration through the disk. However, the precise measurements of HD 209458b suggests that at least some hot Jupiters form beyond the snow line and migrate without accreting significant oxygen-rich solids. As the measurement precision improves, it is our prediction that most if not all planet atmosphere measurements will show super-stellar C/O ratios and/or substellar O/H ratios. The James Webb Space Telescope (JWST) will be able to obtain much more detailed spectra, with 100 ≤ λ/δλ < 3600 between 1 and 11 μ (Greene et al. 2016), of hot Jupiter atmospheres and allow us to test this hypothesis.

As discussed above, the amount of enhancement can also tell us about the migration history of the planets. If migration begins before the dissipation of the gas disk, then the hot...
Jupiter will accrete gas over a wide range of semimajor axes. This will dilute the signal, perhaps erasing it altogether. The fact that we see a clear signal in HD 209458b tells us that at least in some cases, the majority of the planet atmosphere is accreted beyond the snow line without much pollution from oxygen-rich planetesimals. Additional hot Jupiters with measured H$_2$O abundances were published (Barstow et al. 2016) during the review process of this paper. They also show a tendency toward subsolar O/H ratios, suggesting disk-free migration.

4.1. Solar Versus Stellar

Most previous studies of hot Jupiter atmospheres have assumed solar abundance patterns and asked whether the C/O ratio was Sun-like (0.5) or super-solar (e.g., ~1). Although these questions are interesting, using the stellar abundance pattern and looking for deviations in the planet atmosphere with respect to its own star can tell us a lot more about the planet formation process in these systems (Teske et al. 2014). Noting that the C/O ratio of WASP-12b is approximately solar tells us little about how it may have formed, although it does tell us that any H$_2$O in its atmosphere should be readily detectable. Realizing that the planet C/O may be almost twice its stellar C/O ratio can help us identify where in the disk it formed.

5. CONCLUSIONS

The hot Jupiter HD 209458b has a C/O ratio 2.2 times that of its host star and a substellar O/H ratio. Although all other planets in this study are consistent with having the same C/O ratios as their host star, the mean planet-to-star ratio shows an enhancement in C/O. This would be a clear signal of disk-free migration from beyond the snow line and rule out in situ formation of hot Jupiters for most stars. Improved precision in measuring the composition of exoplanet atmospheres with instruments such as JWST will be able to resolve this issue. Stellar hosts of hot Jupiters are preferentially oxygen-rich and standard core accretion models would enhance the oxygen abundance in the planet atmosphere. A deficit of oxygen would signal low levels of planetesimal accretion or disk-free migration. C/O ratios in planet atmospheres show a larger range than those of host stars. This may be telling us something fundamental about the formation processes of the planets.

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