DENSE MOLECULAR GAS EXCITATION IN NUCLEAR STARBURSTS AT HIGH REDSHIFT: HCN, HNC, AND HCO$^+(J = 6 \rightarrow 5)$ EMISSION IN THE $z = 3.91$ QUASAR HOST OF APM 08279+5255

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ABSTRACT

We report the detection of surprisingly strong HCN($J = 6 \rightarrow 5$), HNC($J = 6 \rightarrow 5$), and HCO$^+(J = 6 \rightarrow 5$) emission in the host galaxy of the $z = 3.91$ quasar APM 08279+5255 through observations with the Combined Array for Research in Millimeter-wave Astronomy. HCN, HNC, and HCO$^+$ are typically used as star formation indicators, tracing dense molecular hydrogen gas [$n$(H$_2$) $> 10^5$ cm$^{-3}$] within star-forming molecular clouds. However, the strength of their respective line emission in the $J = 6 \rightarrow 5$ transitions in APM 08279+5255 is extremely high, suggesting that they are excited by another mechanism besides collisions in the dense molecular gas phase alone. We derive $J = 6 \rightarrow 5$ line luminosities of $L_{\text{HCN}}' = (4.9 \pm 0.6)$, $L_{\text{HNC}}' = (2.4 \pm 0.7)$, and $L_{\text{HCO}^+} = (3.0 \pm 0.6) \times 10^{40} \mu_{\text{L}}$ K km s$^{-1}$ pc$^2$ (where $\mu_{\text{L}}$ is the lensing magnification factor), corresponding to $L'$ ratios of $0.23-0.46$ relative to CO($J = 1 \rightarrow 0$). Such high line ratios would be unusual even in the respective ground-state ($J = 1 \rightarrow 0$) transitions, and indicate exceptional, collisionally and radiatively driven excitation conditions in the dense, star-forming molecular gas in APM 08279+5255. Through an expansion of our previous modeling of the HCN line excitation in this source, we show that the high rotational line fluxes are caused by substantial infrared pumping at moderate opacities in a $\sim 220$ K warm gas and dust component. This implies that standard $M_{\text{dense}}/L'$ conversion factors would substantially overpredict the dense molecular gas mass $M_{\text{dense}}$. We also find a HCN($J = 6 \rightarrow 5$)/HCN($J = 5 \rightarrow 4$) $L'$ ratio greater than 1 ($1.36 \pm 0.31$)—however, our models show that the excitation is likely not “super-thermal,” but that the high line ratio is due to a rising optical depth between both transitions. These findings are consistent with the picture that the bulk of the gas and dust in this source is situated in a compact, nuclear starburst, where both the highly active galactic nucleus and star formation contribute to the heating.

Key words: cosmology; observations – galaxies: active – galaxies: formation – galaxies: high-redshift – galaxies: starburst – radio lines: galaxies

Online-only material: color figures

1. INTRODUCTION

Over the past decade, great progress has been made in understanding the conditions for star formation in gas-rich galaxies out to the highest redshifts. Molecular gas, the prospective fuel for star formation, is now detected in more than 70 galaxies at $z > 1$, allowing us to compare the properties of star-forming environments among different galaxy populations in the early universe (see Solomon & Vanden Bout 2005 for a review). These detections are almost exclusively being obtained in rotational transitions of CO, which (due to its relatively low critical density of $n_{\text{crit}}$(H$_2$) $\sim 300$ cm$^{-3}$) is a good proxy for the total amount of molecular gas in a galaxy.

More focused studies of the dense molecular gas found in the star-forming cores of molecular clouds typically employ observations of high dipole moment, high critical density ($n_{\text{crit}}$(H$_2$) $> 10^4$ cm$^{-3}$) molecules such as HCN, HCO$^+$, and HNC, both in nearby galaxies and out to high $z$ (e.g., Gao & Solomon 2004; Riechers et al. 2006a, 2007a; Gao et al. 2007; Baan et al. 2008; Gracia-Carpio et al. 2008).

We here aim to study, for the first time, the dense, star-forming molecular gas excitation in a high-$z$ galaxy. Constraints on the excitation of gas at very high densities and the physical mechanisms (collisions versus other channels) responsible are crucial to understand in more detail how excitation may influence scaling relations between the dense gas content and star formation rate of galaxies back to early cosmic times. To disentangle excitation effects from other phenomena (such as the chemical composition of the gas), it is crucial to target multiple diagnostic lines in a well-studied, key system. The target of this study is the extremely luminous $z = 3.91$ quasar APM 08279+5255, which has been studied comprehensively in CO and other diagnostics (see, e.g., Weiß et al. 2007; Riechers et al. 2009 for details). In particular, it is one of only two high-$z$ galaxies in which multiple dense molecular gas tracers were detected to date (the other being the Cloverleaf quasar at $z = 2.56$; Barvainis et al. 1997; Solomon et al. 2003; Wagg et al. 2005; Riechers et al. 2006a, 2007b; Garcia-Burillo et al. 2006; Guelin et al. 2007).

In this paper, we report the detection of HCN($J = 6 \rightarrow 5$), HCO$^+(J = 6 \rightarrow 5$), and HNC($J = 6 \rightarrow 5$) emission in the quasar host galaxy of APM 08279+5255 ($z = 3.91$), using the Combined Array for Research in Millimeter-wave Astronomy (CARMA). These observations represent the first extragalactic detections of such high-$J$ lines of the dense gas tracers HCN, HCO$^+$, and HNC, and significantly constrain the physical properties of the dense gas in the star-forming regions of this distant galaxy. We use a concordance, flat $\Lambda$CDM cosmology.

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throughout, with $H_0 = 71 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_M = 0.27$, and $\Omega_{\Lambda} = 0.73$ (Spergel et al. 2003, 2007).

2. OBSERVATIONS

We used CARMA to observe the HCN($J = 6\rightarrow5$) ($\nu_{\text{rest}} = 531.7164$ GHz), HCO*($J = 6\rightarrow5$) (535.0618 GHz), and HNC($J = 6\rightarrow5$) (543.8974 GHz) transition lines (~560 $\mu$m) toward APM 08279+5255. At $z = 3.911$, these lines are redshifted to 108.270, 108.952, and 110.751 GHz (~2.8 mm). The target was observed with 14 or 15 antennas (corresponding to 91 or 105 baselines per antenna configuration) for 14 tracks in 2008 March and June (setup 1; HCN/HCO*, observed simultaneously) and 23 tracks between 2009 February and July (setup 2; HNC), amounting to a total observing time of 148 hr. This results in 37 hr on source time for HCN/HCO* and 52 hr for HNC. All HCN/HCO* observations were carried out in D array (11–148 m baselines), while HNC observations were carried out in the C, D, and E arrays (6–185 m baselines after flagging). Observations before 2008 June were carried out with the previous generation 3 mm receivers, and observations after 2008 March were carried out with the new generation 3 mm receivers (which offer improved noise temperatures, tuning range, and stability).

Weather conditions scaled between acceptable and excellent for observations at 3 mm wavelengths. For the HCN/HCO* observations, typical median phase rms values were 275 $\mu$m (normalized to a 100 m baseline, measured at 45° elevation), median optical depths were $\tau_{230\text{GHz}} = 0.24$, and median precipitable water vapor columns were 3.9 mm. For the HNC observations, typical median phase rms values were 300 (C array), 235 (D array), and 395 $\mu$m (E array), median optical depths were $\tau_{230\text{GHz}} = 0.38$, and median precipitable water vapor columns were 5.5 mm. The nearby source J0818+423 (distance to APM 08279+5255: 10:6) was observed every 20 minutes for secondary amplitude and phase calibration. The strong calibration sources J0423−013, J0927+390, 3C 111, 3C 84, and 3C 273 were observed at least once per track for bandpass and secondary flux calibration. Absolute fluxes were bootstrapped relative to Mars, Uranus, or 3C 84 (when no planet was available). Pointing was performed at least every 2–4 hr on nearby sources, using both radio and optical modes. The resulting total calibration is estimated to be accurate within ~15% (see the Appendix).

The 3 mm receivers were tuned between the HCN and HCO* lines at 108.611 GHz (setup 1) and at the redshifted HNC frequency of 110.751 GHz (setup 2; HNC), covering all 10 lines (HCN, HNC, and HCO+) at the redshifted frequency of 108.270 GHz (~2.8 mm). The line wings were resolved at the resolution of our observations. The final rms noise in the LSB is by 3% (HCN/HCO*) and 15% (HNC) better compared to the USB due to atmospheric transmission.

3. RESULTS

3.1. Emission Lines of High-density Gas Tracers

We have detected emission from the $J = 6\rightarrow5$ HCN, HCO*, and HNC emission lines toward the $z = 3.91$ quasar APM 08279+5255. The left panel of Figure 1 shows the HCN and HCO* spectrum at 86 km s$^{-1}$ (31.25 MHz) resolution. Both lines are detected simultaneously on top of strong 2.76 mm continuum emission. The middle panel of Figure 1 shows the HNC spectrum at similar (85 km s$^{-1}$) resolution. The line is also detected on top of strong continuum emission at similar wavelength (2.71 mm). Figure 2 shows the HCN and HCO* velocity channel maps at the same velocity resolution and coverage as the spectrum in Figure 1. At an rms of 0.95 mJy beam$^{-1}$, APM 08279+5255 is detected at 8σ and 6σ significance in the peak channels of the HCN and HCO* lines. Both emission lines are detected over 10 channels on top of the continuum, showing a decline in strength toward the line wings, as expected. Also, marginal excess flux is detected at the peak position of the redshifted HCN lines ($J = 59\rightarrow58$), which is not formally detected. Figure 3 shows the HCN and HCO* velocity channel maps at the same velocity resolution and coverage as the spectrum in Figure 1 (central 12 channels). At an rms of 0.88 mJy beam$^{-1}$, the source is detected at 6σ significance in the peak channel of the HCN line, also showing a decline in flux toward the line wings.

From simultaneous Gaussian fitting to the profiles of the HCN($J = 6\rightarrow5$) and HCO*(J = 6→5) lines and the underlying continuum emission, we derive HCN and HCO* line peak flux densities of 4.04 ± 0.61 and 2.48 ± 0.56 mJy and FWHM velocities of 385 ± 36 and 385 ± 60 km s$^{-1}$, respectively (in good agreement with the 400 ± 40 km s$^{-1}$ FWHM of HCN $J = 5\rightarrow4$; Weiß et al. 2007). This corresponds to integrated HCN and HCO* line fluxes of 1.65 ± 0.19 and 1.01 ± 0.19 Jy km s$^{-1}$. A simultaneous fit to the HCN($J = 6\rightarrow5$) line and underlying continuum emission yields a HCN line peak flux density of 1.57 ± 0.51 mJy and a line FWHM of 520 ± 160 km s$^{-1}$ (also consistent with the FWHM of HCN $J = 5\rightarrow4$ within the errors). This corresponds to an integrated HCN line flux of 0.86 ± 0.24 Jy km s$^{-1}$. From the line peak velocities, we derive formal line redshifts of $z = 3.9126 ± 0.0011$ for HCN/HCO* (simultaneous fit), and $z = 3.9144 ± 0.0011$ for HNC, which are consistent with those derived from other molecular emission lines within the errors (e.g., Riechers et al. 2006b; Weiß et al. 2007).
see close to the peak position of the redshifted HC 3N (shown in the right panel of Figure 1). Marginal excess flux is covered spectral range, including the LSBs (combined spectrum HC3N(60→59), as well as SiO(12→11) emission within the covered spectral range, including the LSBs (combined spectrum shown in the right panel of Figure 1). Marginal excess flux is seen close to the peak position of the redshifted HC3N(J = 57→56) line in the overlap region of both LSB setups, which we however consider not detected in the following. We place 3σ limits of 0.34, 0.52, 0.49, and 0.44 Jy km s⁻¹ on the integrated fluxes from these lines, extracted over a fixed velocity range of 400 km s⁻¹ (see Table 1). We also attempted to stack all three HC3N lines, which results in no clear signal above the formal 3σ limit of 0.26 Jy km s⁻¹. We however note that due to the uncertainties in extraction, this “stacked” limit has to be treated with caution.

3.2. Millimeter Continuum Emission

As mentioned above, we have detected ~2.8 mm (rest-frame ~560 μm) continuum emission toward the host galaxy of the z = 3.91 quasar APM 08279+5255. Emission was detected at high signal-to-noise in each sideband of the two frequency setups. Thus, continuum fluxes were extracted by fitting a two-dimensional, elliptical Gaussian to the source in the u−v plane for each sideband and frequency setting, excluding ranges where line emission was detected. The individual values are listed in Table 2, and are fully consistent with the continuum fluxes obtained from the simultaneous line/continuum fits to the spectra as outlined above. A combination of all measurements yields an average continuum flux of 2.08 ± 0.07 mJy at 2.79 mm, consistent with the spectral energy distribution (SED) of the source (Riechers et al. 2009).

3.3. Line Luminosities and Ratios

From the line intensities, we derive line luminosities and limits of $L'_{\text{HCN}(6\rightarrow5)} = (4.9 \pm 0.6)$, $L'_{\text{HCN}(6\rightarrow5)} = (3.0 \pm 0.6)$, $L'_{\text{HNC}(6\rightarrow5)} = (2.4 \pm 0.7)$, $L'_{\text{HCN}(57\rightarrow56)} < 1.0$, $L'_{\text{HNC}(59\rightarrow58)} < 1.5$, $L'_{\text{HNC}(60\rightarrow59)} < 1.4$, and $L'_{\text{SiO}(12\rightarrow11)} < 1.3 \times 10^{10} \mu L^{-1}$ K m s⁻¹ pc² (Table 1; not corrected for the lensing magnification factor of $\mu L = 4.2$; Riechers et al. 2009).

This corresponds to $J = 6\rightarrow5$ $L'$ ratios of HCN/CO = 0.46 ± 0.07, HCO⁺/CO = 0.28 ± 0.06, and HNC/CO = 0.23 ± 0.07 (relative to CO $J = 1\rightarrow0$; Riechers et al. 2009). Such high line ratios would be unusual even in the respective ground-state ($J = 0$).
Figure 2. Channel maps of the data shown as a spectrum in Figure 1 (left) at the same velocity resolution. Frequencies increase with channel number from left to right, i.e., red to blue. The peak channels of the redshifted HCN, HCO\(^+\), and HC\(_3\)N lines are indicated. Contours are shown at \((-3, -2, 2, 4, 5, 6, 7, 8)\) \(\sigma\) (1\(\sigma\) = 0.95 mJy beam\(^{-1}\)). The beam size (4.5\(\arcsec\) × 3.6\(\arcsec\)) is shown in the bottom left corner. The cross indicates a position of 08\(^{h}\)31\(^{m}\)41\(^{s}\)0, +52\(^{\circ}\)45\(^{\prime}\)17\(^{\prime\prime}\).1 between the CO lens images (Riechers et al. 2009).

We also find a HCN/HCO\(^+\) \(J = 6\rightarrow 5\) \(L'\) ratio of 1.7 ± 0.4, which is higher than the HCN/HCO\(^+\) \(J = 5\rightarrow 4\) \(L'\) ratio of 1.0 ± 0.2 (Weiß et al. 2007; García-Burillo et al. 2006). Such a steeply rising HCN/HCO\(^+\) line ratio with \(J\) is not expected, as the densities and temperatures to collisionally excite HCN and HCO\(^+\) are similar. We note that the \(J = 5\rightarrow 4\) lines were not observed simultaneously, so the error bars of their ratio may be underestimated due to the uncertainties in their relative flux scales. In addition, the difference in line ratio with \(J\) may indicate that HC\(_3\)N(\(J = 49\rightarrow 48\)) line emission contributes significantly to the HCO\(^+\)(\(J = 5\rightarrow 4\)) flux reported by García-Burillo et al., as these lines are blended. The HC\(_3\)N limits derived above would be consistent with up to a 30%-40% contamination (extrapolated from these higher-\(J\) transitions) of the HCO\(^+\)(\(J = 1\rightarrow 0\)) transitions, and indicate exceptional dense gas excitation conditions in this source.
5 → 4) flux. Note that such high HCN/HCO⁺ line ratios are actually observed in nearby infrared-luminous galaxies (albeit in lower-J lines; e.g., Aalto et al. 2007a).

Furthermore, we find HNC(J = 6 → 5)/HCN(J = 6 → 5) = 0.50 ± 0.15 and HNC(J = 6 → 5)/HCO⁺(J = 6 → 5) = 0.82 ± 0.27. The HNC(J = 5 → 4) line was observed toward APM 08279+5255, but it is strongly blended with CN(ν = 4 → 3) (Guelin et al. 2007). The single Gaussian fit by Guelin et al. to the blend of these lines can be translated to upper limits of HNC(J = 5 → 4)/HCN(J = 5 → 4) ≤ 1.12 and HNC(J = 5 → 4)/HCO⁺(J = 5 → 4) ≤ 1.15. Guelin et al. estimate that the CN line contributes ~1/3 to the total flux they measure. Assuming their decomposition is correct, we find HNC(J = 5 → 4)/HCN(J = 5 → 4) = 0.64 ± 0.32 and HNC(J = 5 → 4)/HCO⁺(J = 5 → 4) = 0.66 ± 0.33. Given the similar excitation densities and temperatures of HNC compared to HCN and HCO⁺, we assume that the J = 5 → 4 and 6 → 5 line ratios should be comparable. This assumption is consistent with the observed HNC/HCN ratios if the CN contribution to the HNC(J = 5 → 4) L′ is as suggested by Guelin et al. It also is consistent with the observed HNC/HCO⁺ ratios, in particular if the HCN contribution to HCO⁺(J = 5 → 4) is comparable to the CN contribution to HNC(J = 5 → 4). However, given the remaining uncertainties, observations of another CN transition would be desirable to test this scenario (see also discussion below).

Interestingly, we find a HCN(J = 6 → 5)/HCN(J = 5 → 4) L′ ratio of 1.36 ± 0.31. A ratio of greater than 1 is not expected for optically thick, collisionally excited emission that is thermalized into account, using the HCN collision rates from Schöier et al. (2005), and including the first rovibrational bending mode at 14.0 μm for pumping (ν2 = 1; e.g., Thorwirth et al. 2003). We adopted a HCN abundance per velocity gradient of [HCN]/(dν/dr) = 1 × 10⁻⁹ pc (km s⁻¹)⁻¹ (Helfer & Blitz 1997; Wang et al. 2004). All parameters are required to also reproduce the CO excitation of the source (Weiß et al. 2007), assuming [HCN]/[CO] = 10⁻⁶, and are required to be consistent with the dust SED (Weiß et al. 2007; Riechers et al. 2009). The best solutions were obtained for a spherical, single-component model with kinetic temperatures of Tₖin = 220 K, gas densities of n_gas = 10⁶ cm⁻³, an infrared radiation field temperature of T_IR = T_dust = Tₖin = 220 K, and infrared filling factors⁸ of IRff = 0.3–0.7 (Figure 4; where solutions with higher IRff prefer slightly lower Tₖin to stay consistent with the observed size of the source in CO emission; Riechers et al. 2009).

Due to the high intensity of the IR-pumping field (as indicated by IRff) in these models, HCN(J = 5 → 4) and HCN(J = 6 → 5) are only moderately optically thick (typical optical depths of τ₅−₄ = 0.25 and τ₆−₅ = 0.4), and the optical depth still rises with J between these lines. This implies that the excitation of HCN is not “super-thermal” (as in a “true” population inversion between energy levels), but that the high HCN L′ ratio is merely an optical depth effect. The HCN excitation is dominated by IR pumping for all models that fit the data. Interestingly, the models imply that the HCN emission cannot dominantly arise from a relatively cold, dense gas, and dust component (see Weiß et al. 2007 for multi-component models). This means that standard conversion factors α_HCN = M_dense/|=HCN from HCN line luminosity to dense gas mass (e.g., Gao & Solomon 2004) will substantially overpredict the dense molecular gas mass in this system. Also, the models do not require unusually high relative HCN abundances, contrary to previous suggestions based on a narrower exploration of the parameter space in excitation conditions (Garcia-Burillo et al. 2006).

4. HCN LINE EXCITATION MODELING

To understand the unusual dense gas excitation in APM 08279+5255 in more detail, we have carried out large velocity gradient (LVG) models of the HCN line excitation (see Weiß et al. 2007 for initial study). Our models take both collisional and radiative excitation via infrared (IR) pumping into account, using the HCN collision rates from Schöier et al. (2005), and including the first rovibrational bending mode at 14.0 μm for pumping (ν2 = 1; e.g., Thorwirth et al. 2003). We adopted a HCN abundance per velocity gradient of [HCN]/(dν/dr) = 1 × 10⁻⁹ pc (km s⁻¹)⁻¹ (Helfer & Blitz 1997; Wang et al. 2004). All parameters are required to also reproduce the CO excitation of the source (Weiß et al. 2007), assuming [HCN]/[CO] = 10⁻⁶, and are required to be consistent with the dust SED (Weiß et al. 2007; Riechers et al. 2009). The best solutions were obtained for a spherical, single-component model with kinetic temperatures of Tₖin = 220 K, gas densities of n_gas = 10⁶ cm⁻³, an infrared radiation field temperature of T_IR = T_dust = Tₖin = 220 K, and infrared filling factors⁸ of IRff = 0.3–0.7 (Figure 4; where solutions with higher IRff prefer slightly lower Tₖin to stay consistent with the observed size of the source in CO emission; Riechers et al. 2009).

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5. DISCUSSION

5.1. HCO⁺ and HNC Excitation

Due to the remaining uncertainties in the HCO⁺(J = 5 → 4) and HNC(J = 5 → 4) line intensities (see discussion above), it

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⁸ IRff describes the solid angle fraction of the gas that is exposed to the IR field, which in turn is described by a greybody spectrum.
is currently not possible to properly constrain similar models for the HCO$^+$ and HNC line excitation. However, due to similar rovibrational bending modes ($v_2 = 1$ at 12.1 and 21.6 $\mu$m), it is likely that similar conclusions hold for the HCO$^+$ and HNC excitation. The bending modes of HCN, HCO$^+$, and HNC all lie close to the peak of the dust SED of APM 08279+5255 (Weiß et al. 2007; Riechers et al. 2009), which would make such a scenario plausible. Also, the organic compound HC$_3$N has several bending modes in this wavelength regime (e.g., Wyrowski et al. 1999), which would make a substantial HC$_3$N ($J=49\rightarrow48$) contribution to HCO$^+$( $J=5\rightarrow4$) plausible. Luminous, high-$J$ HC$_3$N emission consistent with radiative excitation scenarios is observed in nearby galaxies like NGC 4418 (Aalto et al. 2007a). On the other hand, the lowest rovibrational transition of the diatomic molecule CN lies at 4.9 $\mu$m (see also Guelin et al. 2007), which lies substantially beyond the peak of the SED. No strong 4.9 $\mu$m (observed frame 24 $\mu$m) CN absorption is detected toward the bright mid-infrared continuum of APM 08279+5255 in deep *Spitzer Space Telescope* spectroscopy (D. Riechers et al. 2010, in preparation). Thus, IR pumping is likely substantially less effective for CN (relative to HCN, HCO$^+$, and HNC), which may contrast with the relatively high CN($N=4\rightarrow3$) luminosity suggested by Guelin et al. (2007), depending on the actual CN abundance. However, luminous CN($N=3\rightarrow2$) emission was detected toward another distant lensed galaxy, the Cloverleaf quasar ($z = 2.56$; Riechers et al. 2007b).

If CN and/or HC$_3$N transitions close to those blended with HNC($J=5\rightarrow4$) and HCO$^+$( $J=5\rightarrow4$) are found to be much fainter than expected based on the $L_\nu$ estimates by Guelin et al.

Figure 4. HCN excitation ladder (spectral line energy distribution; points) and LVG models (lines) for APM 08279+5255, accounting for both collisional and radiative excitation. The HCN($J=5\rightarrow4$) data point is from Weiß et al. (2007). The models give $n_{\text{gas}} = 10^{-2}$ cm$^{-3}$, $T_{\text{kin}} = T_{\text{gas}} = 220$ K, and infrared field-filling factors of $f_{\text{IR}} = 0.3$–0.7 (shown in steps of 0.1).

(A color version of this figure is available in the online journal.)

Recent studies of HCN and HNC emission in nearby infrared-luminous galaxies have revealed sources with high HCN/HNC ratios and high HNC excitation. Two scenarios were brought forward to explain these high ratios: IR pumping and/or increased abundances of HNC in the presence of X-ray-dominated regions (XDRs, often found in regions impacted by emission from active galactic nuclei; Aalto et al. 2007b). On the one hand, we find that HCN($J=6\rightarrow5$)/HCO$^+$( $J=6\rightarrow5$) $> 1$, which is the opposite to what is expected in the XDR scenario (Aalto et al. 2007b). However, we note that our LVG models show that the HCN($J=6\rightarrow5$) line (and thus, likely also the HNC $J=6\rightarrow5$ line) is only moderately optically thick, abundance effects thus cannot be ruled out. On the other hand, the HNC($J=6\rightarrow5$)/HCN($J=6\rightarrow5$) ratio of 0.50 $\pm$ 0.15 in APM 08279+5255 also fits well into the IR-pumping scenario: in nearby infrared-luminous galaxies, the gas is warm enough to efficiently pump HNC (through the 21.6 $\mu$m bending mode with an energy level of $h\nu/k = 669$ K and Einstein $A$ coefficient of $A_{\text{IR}} = 10.2$ s$^{-1}$), but not HCN (through the 14.0 $\mu$m bending mode $h\nu/k = 1027$ K and $A_{\text{IR}} = 1.7$ s$^{-1}$; Aalto et al. 2007b), leading to a high HNC/HCN $L_\nu$ ratio in high-$J$ transitions. In APM 08279+5255, the dust and gas are warm enough to also efficiently pump HCN at high rates, so the line ratio depends mostly on the fraction of the HCN or HNC-emitting gas that is exposed to the IR radiation field and optical depth effects (given that the fraction of $L_\nu$ in the $J = 6\rightarrow5$ transitions due to collisional excitation is likely small). This allows for a broad range of HCN/HCN ratios, and thus is consistent with the comparatively low $J = 6\rightarrow5$ ratio, especially in combination with the high excitation of both molecular line spectral energy distributions (SLEDs). Similar arguments can be made for the relative strength of HCO$^+$, which has its fundamental 12.1 $\mu$m bending mode at slightly higher energy than HCN.

5.2. HNC/HCN Ratio

Given the SED and brightness of the high-$J$ rotational lines of HCN, HNC, and HCO$^+$, pumping by mid-IR rovibrational transitions appears plausible. As the abundance of these molecules is high enough to yield bright rotational lines in emission, the question arises whether or not the pumping transitions themselves may be detectable (which would place better constraints on the abundance of these molecules). The 14.0 $\mu$m HCN feature was detected in absorption in nearby (U)LIRGs (Lahuis et al. 2007). Indeed, sources such as Arp 220 or NGC 4418 show deep absorption features (10%–30% of the continuum flux), indicating high HCN abundances, and evidence for pumping of the rotational HCN lines. The $v_2 = 1$ lines of HCO$^+$, HCN, and HNC are at rest-frame 12.1, 14.0, and 21.6 $\mu$m, i.e., observed frame 59.4, 68.8, and 106.1 $\mu$m, which is within the wavelength range of the Photodetector Array Camera and Spectrometer (PACS) on board the *Herschel Space Observatory* (e.g., Poglitsch et al. 2010). APM 08279+5255 has continuum fluxes of 511 $\pm$ 51 and 951 $\pm$ 228 mJy at 60 and 100 $\mu$m (Irwin et al. 1998). Thus,
absorption features that have a depth of 10% of the continuum flux would be detectable within a few hours with Herschel.

6. CONCLUSIONS

Based on CARMA observations of HCN, HNC, and HCO\(^+\)(J = 6 → 5) toward the z = 3.91 quasar APM 08279+5255, evidence is consolidating that the dense, star-forming molecular gas in this source is substantially enhanced in brightness by IR pumping (on top of gravitational magnification of the source).\(^\text{10}\) The unusual, hundreds of parsec scale, warm gas, and dust in this source appears to harbor a strong IR-radiation field that efficiently pumps the high-J transitions of HCN, HNC, and HCO\(^+\) at rates well beyond those achieved by collisional excitation alone. At least in the case of HCN, this leads to moderately optically thick emission, which appears as “super-thermal” between J = 5 → 4 and 6 → 5 (due to a combination of high excitation and rising optical depth). These findings are consistent with the picture that the bulk of the gas and dust in this source is situated in a compact, nuclear starburst, where both the highly active galactic nucleus and star formation contribute to the heating.

Even though APM 08279+5255 is known to be an extreme system, radiative excitation of dense molecular gas tracers may also play a significant role in other high redshift galaxies, as it does in nearby infrared-luminous galaxies (e.g., Aalto et al. 2007b). This issue becomes increasingly important for the interpretation of observations of higher-J (J > 2) transitions of dense gas tracers such as HCN, HCO\(^+\), and HNC, which are redshifted into the millimeter observing windows at high z, and thus will be primarily targeted toward galaxies in the early universe by future facilities such as the Atacama Large Millimeter/Submillimeter Array (ALMA).

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APPENDIX

FLUX CALIBRATION

The observations presented here were taken over the course of four (HCN & HCO\(^+\))/six (HNC) months. It thus is necessary to correct the flux scale of individual tracks for the intrinsic variability of the phase calibrator. Given the relatively large number of tracks (14/23), this leads to a substantially improved flux calibration relative to the sparsely sampled, internal look-up table of MIRIAD.

Figure 5 shows the variability of J0818+423, the phase calibrator, bootstrapped relative to Mars, Uranus, and 3C84 (top panel), measured over the first band in the LSB at \(\sim 105.3\) GHz. For illustration, a fit to the data is shown. Fluxes for tracks that were observed without a primary calibrator were estimated based on a weighted average of those of tracks that are adjacent in time. As 3C84 is a millimeter-variable source itself,\(^\text{11}\) the bottom panel shows its variability over the same time range in which this project was observed. Within the calibration errors, the phase calibrator shows a clear excess variability over that of 3C84 in the 2009 measurements (data points at (JD=2,454,500)>332). Also, the scatter of values in the 2009 measurements bootstrapped relative to the Mars planetary model is substantially larger than that in 2008 (typically observed under comparable weather conditions). Together with the fact that the fluxes bootstrapped relative to 3C84 agree well with those bootstrapped relative to planet models, this suggests that J0818+423 shows significant intrinsic flux variations over the observed time range. Thus, we obtain the best calibration by

\[ \text{\textsuperscript{10}} IR \text{ pumping may also explain the unusual excitation of low-ionization rest-frame ultraviolet absorption lines in proximate absorbers along selected sight lines toward the active galactic nucleus (which may arise from diffuse gas in the galaxy; Srianand & Petitjean 2000).} \]

\[ \text{\textsuperscript{11}} \text{See, e.g., SMA data archive; http://sma1.sma.hawaii.edu/callist/callist.html.} \]
using and interpolating the individually bootstrapped fluxes for each track, rather than using an averaged value for J0818+423. We conservatively estimate that this flux calibration leads to an overall accuracy of $\sim 15\%$.

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