ABSTRACT

R-band photometric and velocity dispersion measurements for a sample of 452 elliptical and S0 galaxies in 28 clusters are used to construct a template $D_{n,\sigma}$ relation. This template relation is constructed by combining the data from the 28 clusters, under the assumption that galaxies in different clusters have similar properties. The photometric and spectroscopic data used consist of new as well as published measurements, converted to a common system, as presented in an accompanying paper. The resulting direct relation, corrected for incompleteness bias, is $\log D_n = 1.203 \log \sigma + 1.406$; the zero point has been defined by requiring distant clusters to be at rest relative to the cosmic microwave background. This zero point is consistent with the value obtained by using the distance to Virgo as determined by the Cepheid period-luminosity relation. This new $D_{n,\sigma}$ relation leads to a peculiar velocity of $-72 \pm 189$ km s$^{-1}$ for the Coma Cluster. The scatter in the distance relation corresponds to a distance error of about 20%, comparable to the values obtained for the fundamental plane relation. Correlations between the scatter and residuals of the $D_{n,\sigma}$ relation with other parameters that characterize the cluster and/or the galaxy stellar population are also analyzed. The direct and inverse relations presented here have been used in recent studies of the peculiar velocity field mapped by the ENEAR all-sky sample.

Key words: cosmology: observations — galaxies: clusters: general — large-scale structure of universe

1. INTRODUCTION

Present-day elliptical galaxies form a remarkably homogeneous class of objects that obey scaling relations involving their structural and dynamical properties. Indeed, elliptical galaxies are known to populate the so-called fundamental plane (FP; Djorgovski & Davis 1987; Dressler et al. 1987), in a three-dimensional space defined by the surface brightness $\mu_r$, effective radius $r_e$, and internal velocity dispersion $\sigma$. Therefore, by choosing an appropriate combination of parameters, a tight relation between distance-dependent and distance-independent quantities can be found. The $D_{n,\sigma}$ relation is such a relation (Dressler et al. 1987), where $D_n$ is the physical scale of the galaxy defined at a specified surface brightness level ($D_n \equiv d_nR_e$, where $d_n$ is a measure of the angular size of the galaxy and $R_e$ is the distance of the galaxy). The existence of such scaling relations provides an important tool for studying the properties of the stellar populations and the evolution of elliptical galaxies (e.g., Jørgensen, Franx, & Kjærgaard 1996; Franx et al. 1997) and for constraining models of spheroidal formation (e.g., Bressan, Chiosi, & Fagotto 1994; Baugh, Cole, & Frenk 1996). Furthermore, these relations provide the means of measuring relative distances to early-type galaxies. This is the primary goal of this work. At the present time, some doubts remain whether these relations depend on the environment (e.g., Gibbons, Fruchter, & Bothun 2001); if they do, this would lead to the measurement of spurious motions. In addition, until recently it was unclear how distances derived using $D_{n,\sigma}$ related to those measured by the Tully-Fisher (TF) relation (Scodellino 1997; Scodellino, Giovanelli, & Haynes 1997). The FP and $D_{n,\sigma}$ scaling relations are not entirely equivalent (Jørgensen, Franx, & Kjærgaard 1993), and the $D_{n,\sigma}$ relation is expected to be less accurate if the range of galaxy sizes is large (Kelson et al. 2000). In addition, it has been claimed that the scatter around the FP is smaller, suggesting that $D_{n,\sigma}$ distances are less accurate (Jørgensen et al. 1993). However, recent studies do not seem to support these claims (Jørgensen et al. 1996; D’Onofrio et al. 1997; Hudson et al. 1997). While the FP relation is usually used for detailed cluster studies, for large samples such as the magnitude-limited, all-sky sample of early-type galaxies (ENEAR; da Costa et al. 2000b) presented below, the number of galaxies with available $d_n$ measurements is ~50% larger than the number with available FP measurements. Therefore, since the primary goal of this project has been to estimate galaxy distances and derive the peculiar velocity field, we have focused our attention on the derivation of a template $D_{n,\sigma}$ relation. A similar analysis for the FP relation will be presented in a future paper.

This work uses the ENEARc sample of early-type galaxies in 28 clusters presented by Bernardi et al. (2002, here-
after B02). This sample combines data available in the literature with new measurements converted into a common system thanks to the effort of securing new measurements for a large number of galaxies in common with previously available samples. Another important feature of the sample is that cluster membership was carried out using groups identified in complete redshift surveys of the nearby universe, thereby leading to a more systematic assignment than was possible in earlier work.

In this paper we obtain $D_n\sigma$ fits for each of the 28 clusters accounting for various possible biases. We then combine the sample to construct a global template relation under the assumption that early-type galaxies in different clusters are similar. We also study the residuals with respect to the template to investigate, a posteriori, the accuracy of this assumption. The resulting relation is used to compute peculiar velocities of clusters as well as of the ENEAR all-sky sample of early-type galaxies. Both these samples have been used to measure the bulk flow velocity, to set constraints on cosmological parameters, and to characterize the velocity field and mass distribution in the local universe (da Costa et al. 2000a; Borgani et al. 2000a; Nusser et al. 2001; Zaroubi et al. 2001).

The paper is organized as follows. In § 2 we describe the data set with the individual galaxy parameters. In § 3 we present the calibration of the direct (forward) $D_n\sigma$ relation. We quantify the selection bias, which, if not corrected for, can lead to an erroneous determination of the distance relation coefficients and its scatter. This is critical for studies of the cosmic flow field and the properties of early-type galaxies. The direct relation has been applied to compute distances for the galaxy sample used in da Costa et al. (2000a) and Zaroubi et al. (2001). Also shown are the parameters for the inverse relation obtained by regressing on the distance-independent quantity $\log \sigma$; this inverse relation has been used in the analysis of the peculiar velocity field of clusters and “field” galaxies in redshift space (Borgani et al. 2000a; Nusser et al. 2001). The measured distances and peculiar velocities for the ENEARc sample are reported in § 4. In § 5 we look for potential systematic effects that may invalidate the underlying assumption that galaxies in different clusters are similar. Finally, in § 6 we present a brief summary of our results.

2. THE CLUSTER SAMPLE

The spectroscopic and photometric parameters for the 452 galaxies in 28 clusters used here are presented in B02, where we describe the selection of the cluster sample and membership assignments. The clusters we consider in the present study span the redshift range $1000 < cz \leq 11,000 \text{ km s}^{-1}$, covering both equatorial hemispheres. The characteristic parameters (mean redshift, size, and velocity dispersion) of nearly all clusters were computed from the analysis of “groups” identified using objective friends-of-friends algorithms applied to complete redshift surveys (see B02). Exceptions include the Centaurus complex, three clusters previously studied by Jørgensen, Franx, & Kjærgaard (1995a, 1995b) and Jørgensen (1997) (A539, AS 639, and A3381), and two observed by Smith et al. (1997) (7S 21 and A347). The parameters adopted for these cases and the reasons for including them are discussed in B02. Using the identified groups as signposts for clusters, galaxies fainter than the ENEAR magnitude limit ($m_B \sim 14.5$) were considered members by adopting well-defined position and kinematic criteria that should minimize errors in the membership assignment. About 2% of the galaxies previously assigned to clusters were found not to be members according to the membership criteria adopted.

The data set of the ENEARc sample is a compilation including new photometric and spectroscopic measurements obtained as part of this program as well as data previously reported in the literature. In B02 we presented the photometric and spectroscopic measurements for 640 individual cluster galaxies, including new measures of the photometric parameter $d_n$ for 348 galaxies and new spectroscopic measurements of redshift, velocity dispersion, and the Mg2 index for 229 galaxies. Our new data for cluster galaxies have been combined with those in the literature by converting all measurements to a common system (see B02). This was possible by securing observations for a representative number of galaxies in common with other samples, thus allowing the definition of conversion relations. Data from the literature come from Dressler (1987), Lucey & Carter (1988), Faber et al. (1989), Dressler, Faber, & Burstein (1991), Jørgensen et al. (1995a, 1995b), Lucey et al. (1997), and Smith et al. (1997). A detailed description of the new R-band imaging data and parameters, including the total magnitude, effective radius $r_e$, mean surface brightness within this radius $\mu_e$, and disk-to-bulge ratio $D/B$, will be presented in Alonso et al. (2002), while the spectroscopic data will be presented by Wegner et al. (2002).

In constructing the $D_n\sigma$ relation, we exclude 188 galaxies (see B02, Table 8) either because they present photometric or spectroscopic features typical of later type galaxies (e.g., presence of arms, bar, dust lane, emission lines) or because the measured parameters could be affected as a result of contamination by nearby galaxies or stars. Pruning the sample in this way decreases the scatter ($\sim 5\%$) but leaves both the slope and zero point essentially unchanged.

The selection and completeness of the sample of early-type galaxies in clusters are not very well defined because (1) the selection process varies from cluster to cluster and (2) each cluster is a compilation of galaxies taken from different sources. The sample of cluster galaxies includes (1) all early-type galaxies brighter than $m_B < 14.5$ (since they were extracted from complete magnitude-limited catalogs; see da Costa et al. 2000b) and (2) fainter early types with photometric and spectroscopic data available in the literature (see B02). Furthermore, for any given signal-to-noise ratio and resolution, there is a lower limit below which the velocity dispersion measurements are unreliable. Since a cluster may have measurements taken from different sources, this lower limit is not well defined either. For example, for some sources in the literature only measurements of $\sigma > 100 \text{ km s}^{-1}$ are available; for our data this limit can be as low as $45 \text{ km s}^{-1}$ owing to the higher resolution used (see Wegner et al. 2002). We have checked that the results presented below are not significantly dependent on the adopted velocity dispersion limit.

3. DETERMINING THE DISTANCE RELATION

3.1. The Method

A galaxy’s angular size varies inversely as $R$, its comoving distance. If $d_n$ is the measured size of the galaxy on the sky, then $1/d_n$ is a measure of its distance. The central velocity
dispersion of a galaxy $\sigma$ is expected to be correlated with its physical size $D_n = d_n R$ (e.g., Dressler et al. 1987). If we measure both $d_n$ and $\sigma$, then the basic distance indicator becomes
\[
\log R = a \log \sigma - \log d_n + b, \tag{1}
\]
where $a$ represents the scaling of velocity dispersion with size, $D_n \propto \sigma^a$. Here both $R$ and $\sigma$ are in units of km s$^{-1}$, and $d_n$ is expressed in units of 0.1. Define the quantities $y \equiv \log D_n + \log h$, where $h = H_0/(100 \text{ km s}^{-1} \text{ Mpc}^{-1})$, and $x \equiv \log \sigma$. Then the distance relation is
\[
y = ax + b. \tag{2}
\]
The slope of equation (2) is usually determined using cluster galaxies because they can be assumed to be equally distant and the uncertainty in the estimated distances falls as $1/N^{1/2}$. If a distance relation that does not depend on cluster properties exists, it can be determined by combining the data from all available clusters to produce a standard template relation. Although the peculiar velocity field is unknown, combining many different clusters should improve the statistical accuracy of the slope and zero point. (Note that here we are only assuming that early-type galaxies in different clusters are similar; we are not addressing the possible differences between galaxies in clusters and in regions of lower density.)

Here the template parameters—zero point, slope, and relative motions of each cluster—are determined simultaneously. Such a procedure has been adopted by a number of authors (Baggley 1996; Giovanelli et al. 1997, hereafter G97; Scodellio 1997; Scodellio et al. 1998; Colless et al. 2001) in determining the $D_n-\sigma$, $FP$, and $TF$ relations. Our notation below follows G97. The distance relation can be derived by either a direct (forward) or an inverse linear regression fit, depending on whether the slope is obtained using the distance-dependent $d_n$ or the distance-independent parameter $\sigma$ as the independent variable. We study the direct first and the inverse later, in § 3.3.

The coefficients of the direct relation $y = a_d x + b_d$ are determined as follows. We have $N_g$ galaxies in $N_{cl}$ clusters. Let $(x_{ik}, y_{ik})$ denote the values of $x$ and $y$ for the $i$th galaxy in the $k$th cluster. For the $k$th cluster, the distance relation is
\[
y = a_k x + b_k. \tag{3}
\]
Our assumption that the distance relation does not depend on the properties of a cluster means that $a_k$ has the same value, $a_{ik}$, for all clusters. If as a first guess one uses the observed radial velocity $cz$ as the “distance” $[y \equiv \log D_n + \log h \equiv \log (d_n cz) + \log h]$, the zero point $b_k$ is different for different clusters only because of their peculiar velocities relative to the Hubble flow: $D_k = b_d + \Delta_k$. We would like to find those values of $a_{ik}, b_{ik}$, and $\Delta_k$ for which the scatter around the mean relation is minimized. Therefore, we minimize
\[
\chi^2 = \sum_{k=1}^{N_{cl}} \sum_{i=1}^{N_k} \left[ y_{ik} - (a_{ik} x_{ik} + b_{ik} + \Delta_k) \right]^2 / \sigma_{ik}^2, \tag{3}
\]
with respect to the slope $a_{ik}$, the zero point $b_{ik}$, and the relative offsets $\Delta_k$. Here $\sigma_{ik}$ is related to the measurement error in $D_n$ of the $i$th galaxy in the $k$th cluster.

If “distant” clusters (which we define as being clusters beyond 3000 km s$^{-1}$) are at rest relative to the cosmic microwave background (CMB), then the sum over their peculiar velocities should equal zero. Therefore, once $a_{ik}, b_{ik}$, and the $\Delta_k$ values have been found, we compute
\[
\frac{\sum_k N_g(k) \Delta_k}{\sum_k N_g(k)},
\]
where the sum is over the subset of “distant” clusters in our sample. We then subtract this value from each of the $\Delta_k$ values. In effect, this sets the overall zero point of the distance relation.

As will be shown below, this condition turns out to be equivalent to (1) requiring that the distance to Virgo be equal to that given by the Cepheid period-luminosity relation (Kelson et al. 2000) or (2) assuming that the Coma Cluster is at rest.

Formally, the equations above fully describe the procedure we use to determine the parameters that describe the distance relation. However, when working with real data, one must also consider possible sources of bias (for a review see Strauss & Willick 1995 and references therein), as described below.

### 3.2. Monte Carlo Bias Correction

For the direct relation, i.e., when fitting on the distance-dependent parameter, the most pernicious bias is that due to incompleteness. This bias leads to a shallower slope, a larger zero point, and an underestimate of the scatter. Although analytic bias correction schemes have been proposed (Willick 1994), the assumptions made are hardly met by real data. This bias is particularly difficult to handle when the completeness varies from cluster to cluster, as is the case in our sample. Here we follow G97, Scodellio (1997), and Scodellio et al. (1998) and use a Monte Carlo approach to estimate the bias correction, although this is not the only method that can be used (Wegner et al. 1996).

As mentioned above, to estimate the bias, we must first know the incompleteness in $D_n$ for each cluster. This requires knowledge of the $D_n$ distribution function, the counterpart of the luminosity function. Since this function is not directly available, two approaches are possible: (1) assume that a fair representation of this distribution is given by that of a nearby cluster that is complete, or (2) examine the correlation of $d_n$ with some other measure of the angular size of a galaxy, whose distribution is known. We adopt the second approach.

Let $\theta_{25}$ denote the angular diameter enclosing an integrated surface brightness of 25 mag arcsec$^{-2}$, and let $D_{25}$ denote the physical size obtained by multiplying this angular size by the distance to the galaxy. The distribution of $D_{25}$ in the ESO-LV catalog (Lauberts & Valentinij 1989) is
\[
\phi(D_{25})dD_{25} \propto \exp \left( -\frac{D_{25}}{D_*} \right) \frac{dD_{25}}{D_*}, \tag{4}
\]
with $D_* = 2610$ km s$^{-1}$ (Sodré & Lahav 1993). There is a tight correlation between $d_n$ and $\theta_{25}$ (e.g., Wegner et al. 1996), so the distribution of $D_n$ in a complete sample should be well approximated by
\[
\Phi(D_n) = \phi(D_{25}) \frac{dD_{25}}{dD_n}. \tag{5}
\]
The results are insensitive to the exact shape of the diameter function (e.g., G97). The ratio of the observed distribution of $D_n$ in a given cluster with the one expected for a complete sample provides an estimate of the completeness. This ratio
depends on $D_n$ differently for each cluster: we call it the completeness $C_k(D_n)$. Note that $C_k(D_n)$ varies between 0 and 1.

Incompleteness leads to a bias in determining the distance indicator coefficients $a_d$ and $b_d$, which we estimate using the following Monte Carlo approach. In the first step, distances to clusters are approximated by using their redshifts, and the $\chi^2$ defined in equation (3) is minimized. This provides initial guesses for the slope, zero point, peculiar velocities, and scatter $\epsilon$ around the mean relation. The scatter $\epsilon$ may change with velocity dispersion, so we actually compute $\epsilon(x)$ in bins of $x$.

For the $i$th galaxy in the $k$th cluster a bias correction $B_{ik}$ is obtained as follows. A Gaussian zero mean variance random number $g$ is generated. This, with the coefficients $a_d$ and $b_d$ and the scatter $\epsilon(x_{ik})$, is used to compute $\langle y_{ik}^g \rangle = a_d x_{ik} + b_d + g \epsilon(x_{ik})$. This represents the value of $y$ that the observed galaxy may have had. If this value of $y$ was too small, the galaxy would not have been observed. The probability that it would have been observed is proportional to the completeness $C_k(\langle y_{ik}^g \rangle)$. Therefore, we generate a random number $u$ that is distributed uniformly between 0 and 1. The number $y_{ik}^g$ is accepted if $u \leq C_k(\langle y_{ik}^g \rangle)$. We repeat this procedure until we have accepted 500 values of $y_{ik}^g$ for each galaxy. The mean of these values $\langle y_{ik}^c \rangle$ reflects the incompleteness of the real sample. It thus allows a direct estimate of the bias:

$$B_{ik} = \langle y_{ik}^c \rangle - (a_d x_{ik} + b_d).$$

This value is used to define corrected values

$$y_{ik}^c = (a_d x_{ik} + b_d) - B_{ik}.$$

These corrected values are inserted into equation (3); minimizing yields new estimates of $a_d$, $b_d$, the $A_k$ values, and the scatter $\epsilon$. The process is repeated until convergence is reached; applied to our data, this happens after about four iterations.

### 3.3. The Template Distance Relation: Fitting Parameters

We apply the above procedure to the cluster sample presented in §2. We start by assuming that the clusters are at rest relative to the Hubble flow and that their distances are given by the mean cluster redshift (see B02). Figure 1 shows the individual uncorrected cluster data at the start of the iterative process. The solid line represents the best fit after minimizing $\chi^2$ (eq. [3]) for the first time. Note that the number of galaxies in each cluster varies dramatically and, for most groups, only the more luminous, high velocity dispersion cluster galaxies are included in the sample. Therefore, if the selection bias correction is not applied, a significant bias exists in the global template constructed using all clusters. The relative offsets between those data points and the distance relation reflect the relative motions of the clusters.

Figure 2 shows the completeness function $C(D_n)$ for each cluster computed from the ratio of the number of objects observed in the cluster to the number predicted by the fitted diameter distribution function (eq. [5]).

In practice, this is done after binning the data in $\Delta y \equiv \Delta \log D_n = 0.2$ bins and then smoothing with a Hanning filter [convolving with a $(0.25, 0.50, 0.25)$ function] to reduce the effects of small number statistics. The solid curve is a fit to the histograms using the function (G97)

$$C(y) = \frac{1}{1 + e^{(\gamma - y)/\eta}}$$

Table 1 gives the parameters $\gamma$ and $\eta$ of the completeness function for each cluster. At the bright end (large values of $y$) the completeness was normalized to unity, based on the fact that in all clusters the brightest galaxies are always included in the cluster sample. Comparisons between the predicted and observed diameter functions for the nearby Virgo Cluster are in good agreement down to small values of $D_n$, indicating that we could have used Virgo to estimate the completeness of the other clusters.

Using this function as input, we estimated the bias correction $B_{ik}$ for the $i$th galaxy in the $k$th cluster. The results after the final iteration are shown in Figure 3. For nearby clusters, such as Virgo and Fornax, the incompleteness bias correction is small, as expected. For more distant clusters the correction can be significant, with $\Delta y \sim 0.1$ (corresponding to $\Delta m \sim 0.5$ mag).

After the iterating, final values for the distance relation coefficients are determined. Applying the condition that “distant” clusters, i.e., clusters beyond 3000 km s$^{-1}$ (with a mean redshift of 6000 km s$^{-1}$), are at rest with respect to the CMB, and excluding clusters with suspiciously large peculiar velocities (see discussion below), we obtain the follow-
The derived zero point is consistent with the value obtained by assuming that Coma is at rest with respect to the CMB (the $D_n-$parameters relation given above leads to a peculiar velocity of approximately $-72 \pm 189$ km s$^{-1}$ for Coma).

The error in the zero point has two sources. The first is related to the scatter in the distance relation and the procedure adopted in the construction of the template relation. This was estimated as follows. We constructed data sets by randomly removing some points and replacing them with others from the observed data set. For each cluster the same fraction of data points was replaced, typically from 5% to 25%. For clusters with few members ($d \leq 10$) for which this was not possible as a result of the small number of cluster members, we left out one or two observations in sequence. We fixed the slope of the $D_n-$parameters relation and derived the zero point from each simulated data set. The random uncertainty in the zero point is given by the standard deviation of a Gaussian fit to the distribution of these zero points. This yields an error of 0.018, corresponding to an error of $\sim 4\%$ in distance. The second contribution to the zero-point error is the uncertainty in the mean velocity of the distant cluster sample used to set the zero point of the relation. This uncertainty is due to the finite number of clusters used to sample
the peculiar velocity field of the clusters. It is also susceptible to cosmic variance. The uncertainty in the mean peculiar velocity of the cluster sample is given by \( \sigma / N^{1/2} \), where \( \sigma \) is the rms of the clusters’ peculiar velocity distribution and \( N \) is the number of clusters. Note, however, that since the clusters in our sample are not randomly distributed, the actual number of clusters with uncorrelated velocities should be smaller than the 28 clusters considered. Here we estimate the error to be \( \sigma / N^{1/2} \), which corresponds to \( 100 / 4500 \approx 2\% \) at the median distance of the clusters in our sample. Adding in quadrature the different contributions to the error budget, we estimate the final error in the zero point to be 0.021, with the main contribution coming from the uncertainties associated with the distance relation.

Figure 4 shows the initial and final estimates of the distance relation together with the distribution of the observed rms scatter (solid line) and the intrinsic scatter (dashed line), as a function of \( \sigma \). The intrinsic scatter was derived by subtracting the measurement uncertainties in quadrature from the rms scatter of the fit; although it increases at low \( \sigma \), its mean value is \( \approx 0.06 \) dex. This scatter may reflect differences in the stellar populations of the cluster member galaxies. This possibility will be discussed in § 5.2. The intrinsic scatter limits the accuracy of the derived distances. The mean of the total scatter, \( \bar{\epsilon} \approx 0.085 \) dex, yields a distance error \( \Delta \approx 20\% \). This is comparable to the errors obtained using FP relations (e.g., Hudson et al. 1997).

We used the Coma Cluster to test if our correction based on Monte Carlo simulations is reliable. For this cluster we extracted subsamples using different magnitude limits and computed the \( D_n \sigma \) relation for each individual subsample.

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**Fig. 2.**—Selection function for each cluster, computed from the ratio of the number of objects observed in the cluster to the number predicted by the fitted diameter distribution function. Solid curves show fits to the histograms of the form given by eq. (8).
as follows: (1) without applying the bias correction described in § 3.2 and (2) correcting the slope for selection effects. We found that imposing a magnitude limit biases the slope to lower values but that the Monte Carlo technique used recovers the correct value of the slope in each subsample. Using galaxies in Coma, we also checked whether adopting different lower limits in velocity dispersion biases the slope of the distance relation. We found that increasing the value of the velocity dispersion cutoff, for instance, from 70 to 100 km s\(^{-1}\), does not significantly affect our results, with the slope varying by less than 2% and the scatter remaining unchanged.

Figure 5 shows the bias-corrected data points for all clusters and the final fit. A number of interesting cases are evident. For instance, HMS 0122+3305, A2199, and Cen 30 exhibit clear evidence of either spatial substructure or distinct galaxy populations, and the galaxies in the cluster AS 714 do not strictly follow the template relation. The individual \(D_n - \sigma\) exhibit a tilt relative to the template relation. We will return to these points below.

To evaluate the robustness of our results, we derived the \(D_n - \sigma\) relation when specific subsamples of galaxies were excluded. The results are summarized in Table 2: column (1) gives the subsample of galaxies removed (A: “peripheral” objects defined in B02; B: clusters whose individual \(D_n - \sigma\) relations differ significantly \([\Delta a \gtrsim 0.2]\) from eq. [9]; and individual clusters); column (2) gives the number of remaining galaxies that were used to compute the \(D_n - \sigma\) relation; and columns (3), (4), and (5) give the slope, the zero point, and the rms scatter of the \(D_n - \sigma\) relation obtained using the number of galaxies given in column (2). Based on these tests, we conclude that the variation of the slope, \(a\), agrees with the formal error computed from the bootstrap resampling \((\sigma_a \sim 0.023)\).
We have also computed the direct relation using orthogonal fits, allowing for errors in both log $D_n$ and log $\sigma$, and the inverse relation, ignoring the bias correction. The results are shown in Figure 6. The corresponding coefficients and scatter in log $D_n$ are given in Table 3. The inverse relation is insensitive to the photometric selection and is, in principle, bias free if no a priori cut is made in that variable (see Strauss & Willick 1995). However, as some data from the literature in our sample are limited to galaxies with $\sigma \gtrsim 100$ km s$^{-1}$, this assumption may not hold, and perhaps in these cases treatment similar to that carried out for the distance-dependent parameter should be considered.

The three fitting relations in Table 3 clearly show that it is crucial to use a self-consistent fitting algorithm, to have a large and homogeneous set of data, and to correct for selection biases. For this analysis we found that even though the slopes and zero points of the direct, bivariate, and inverse relations differ by more than 2 $\sigma$, the distances of the 28 clusters in the ENEARc sample agree well (see § 4).

The coefficients determined from our sample are compared with those found by other authors in Table 4. Our results are generally in good agreement with previous determinations, except for those of Baggley (1996) and Lucey et al. (1997). Saglia et al. (2001) have recently revised Baggley’s result giving both a slope and zero point comparable to our values. Lucey et al. (1997) used two distant clusters, A2199 and A2634, to obtain their results. Our analysis shows that A2199 has an individual $D_n-\sigma$ relation that differs significantly from equation (9) (see Table 6 below), while A2634 has a high peculiar motion (see § 4 for more details).

4. CLUSTER PECULIAR VELOCITIES

We compute distances to galaxies in clusters using the “direct” template relation found in the previous section. Figure 7 shows the differences in distance between each individual galaxy and its cluster, the distance of which was com-

![Figure 4](http://example.com/figure4.png)
computed as the error-weighted mean of the galaxy distances in the cluster. For the best-sampled clusters (e.g., Virgo, Fornax) the distance distributions have well-defined peaks and small scatter, resulting in good mean distances. However, there are a few complex cases in which clusters exhibit substructure (e.g., HMS 0122+3305, Perseus, Coma, and Cen 30). Furthermore, there are clusters that show large scatter, or poorly defined peaks, or both (e.g., A2199, A2634, and Klem 44). Finally, there are clusters that have only a few galaxies: typically, these are either nearby small groups (e.g., 7S 21, A347, A1367, HG 50, Pegasus, Doradus, AS 714) or very distant clusters (e.g., A3381) with large distance uncertainties.

The cluster distances were corrected for homogeneous Malmquist bias [following Lynden-Bell et al. 1988, the estimated distance is multiplied by exp(3.5e^2/N_g), where N_g is the number of cluster galaxies]; this correction generally amounts to less than ~3% of the distance for the smallest groups.

The radial component of the peculiar velocity of each cluster, \( v_p = c v_{\text{cor}} - R \), was computed using the Malmquist-corrected distance \( R \) and the mean cluster radial velocity \( c z \) presented in B02 and corrected for the cosmological effect:

\[
    c z_{\text{cor}} = c z - \log \left( \frac{1 + (7/4)(cz/c)}{1 + (7/4)(c z_{\text{Coma}}/c)} \right),
\]

where \( c z_{\text{Coma}} \) is the Coma Cluster radial velocity and \( c \) is the speed of light (Lynden-Bell et al. 1988). Baggley (1996) computed a more accurate cosmological correction, but for nearby galaxies equation (10) is a good approximation.

The measured cluster distances and peculiar velocities are presented in Table 5: column (1) gives the cluster name; column (2) gives the number of observed cluster galaxies; col-
Columns (3) and (4) are the cluster’s Galactic coordinates; column (5) is its redshift determined from the group-finding algorithm (see B02) to complete, magnitude-limited redshift surveys (see §2). The figure shows that the fraction of galaxies in some clusters identified by the algorithm is significantly smaller than the total number of galaxies used in this paper (e.g., Coma). For such clusters, using all galaxies (open histograms) or using only this subset (filled histograms) may provide different estimates of the cluster’s mean redshift, velocity dispersion, and other parameters. Nevertheless, the figure suggests that the mean redshift remains about the same, even though the filled histograms are likely to underestimate the dispersion around the mean redshift.

The error-weighted mean cluster redshift of early-type galaxies only (long vertical line) and that given by the group-finding algorithm (short vertical line) are also shown in Figure 8. For the latter, the sample of galaxies assigned to a group/cluster included all morphological types. Note the significant redshift differences, occasionally as large as 300 km s$^{-1}$. The most deviant cases ($\gtrsim 2\sigma$, where $\sigma$ is the error in the mean cluster redshift of early-type galaxies, i.e., the error on the position of the long vertical line) are A347, A539, A1367, Eridanus, Doradus, and Pavo II. This suggests that using a subsample of galaxies in a cluster (especially when only a few objects are selected) to compute the cluster redshift may introduce an error that can, in some cases, be large. This possibility has been ignored in the past and may account for some disparities in the measurements of the peculiar velocity.

The upper panel of Figure 9 shows the distribution of cluster peculiar velocities. The lower panel shows that velocities do not depend on the estimated distances; large peculiar velocities occur at both small and large distances. Filled circles are for the “distant” clusters used in the final calibration of the $D_\sigma$-$\sigma$ relation—the subsample that is required to be at rest relative to the CMB. Open circles indicate nearby clusters plus three additional clusters in the Great Attractor (e.g., Lynden-Bell et al. 1988) region: Cen 30, AS 714, and AS 753. Triangles show clusters with data exclusively from the literature. The 1 $\sigma$ error bars were computed by adding the distance and the cluster mean redshift errors in quadrature. The errors in distances were taken to be $\Delta_0/N^{1/2}$, where $\Delta$ is the fractional distance error derived from

### Table 2

| Objects Removed  | $N_{\text{remain}}$ | $a$  | $b$  | $\epsilon$ |
|------------------|---------------------|------|------|------------|
| A…………………... | 360                 | 1.197| 1.423| 0.084      |
| B…………………... | 374                 | 1.184| 1.448| 0.086      |
| 7S 21……………... | 447                 | 1.204| 1.405| 0.083      |
| Pisces……………...| 433                 | 1.207| 1.407| 0.084      |
| HMS 0122+3305……..| 444                 | 1.199| 1.413| 0.083      |
| A262………………...| 446                 | 1.204| 1.406| 0.083      |
| A347………………...| 447                 | 1.203| 1.407| 0.083      |
| Perseus……………...| 428                 | 1.202| 1.411| 0.084      |
| A539………………...| 440                 | 1.203| 1.406| 0.083      |
| A1307………………..| 448                 | 1.203| 1.407| 0.083      |
| Virgo………………...| 410                 | 1.203| 1.408| 0.079      |
| Coma………………... | 374                 | 1.187| 1.438| 0.086      |
| HG 50………………...| 447                 | 1.203| 1.405| 0.083      |
| A2199………………...| 441                 | 1.194| 1.424| 0.081      |
| Pegasus……………...| 450                 | 1.204| 1.405| 0.083      |
| A2634………………...| 442                 | 1.202| 1.408| 0.083      |
| A194………………... | 439                 | 1.206| 1.400| 0.083      |
| Fornax………………. | 436                 | 1.206| 1.401| 0.082      |
| Eridanus……………...| 441                 | 1.209| 1.397| 0.082      |
| Doradus……………...| 450                 | 1.202| 1.408| 0.083      |
| A3381………………...| 448                 | 1.202| 1.408| 0.083      |
| Hydra………………... | 415                 | 1.204| 1.404| 0.081      |
| AS 639………………...| 448                 | 1.203| 1.407| 0.083      |
| Cen 45………………...| 446                 | 1.203| 1.408| 0.083      |
| Cen 30………………...| 433                 | 1.183| 1.448| 0.080      |
| AS 714………………...| 447                 | 1.203| 1.407| 0.083      |
| Klem 27………………..| 444                 | 1.203| 1.408| 0.083      |
| AS 753………………...| 436                 | 1.205| 1.403| 0.082      |
| Pavo II………………..| 442                 | 1.201| 1.412| 0.083      |
| Klem 44………………..| 446                 | 1.204| 1.406| 0.082      |

Note.—A: peripheral cluster galaxies; B: clusters whose individual $D_\sigma$-$\sigma$ relations differ significantly ([$D_\sigma$ $\geq$ 0.2]) from eq. (9); HMS 0122+3305, A1367, HG 50, A2199, Doradus, A3381, AS 639, Cen 30, and AS 714.

### Table 3

| Type (1)          | $a$  | $b$  | $\epsilon$ |
|-------------------|------|------|------------|
| Direct………………...| 1.203| 1.406| 0.085      |
| Direct orthogonal fit | 1.414| 0.925| 0.089      |
| Inverse………………...| 1.460| 0.826| 0.075      |

Note.—The uncertainty in the distances determined using the inverse relation is $a \times \epsilon$. 

### Table 4

| Source  | Type   | $a$  | $b$  | $\epsilon$ |
|---------|--------|------|------|------------|
| LC88……...| Direct | 1.200| –1.679| 0.090      |
| 7S……...| Direct | 1.200| 1.411| 0.090      |
| DFB91…..| Direct | 1.330| –1.967| 0.110      |
| B96……...| Direct | 0.938| ±0.072| 0.071      |
| JFK96…..| Orthogonal | 1.320| ±0.070| 0.088     |
| L97……...| Direct | 0.913| ±0.090| 0.075      |
| D97……...| Direct | 1.240| ±0.060| 0.080      |
| H97……...| Inverse | 1.419| ±0.044| 0.065      |
| GFB01……...| Inverse | 1.420| ±0.040| 0.059      |

Note a — LC88: Lucey & Carter 1988; 7S: Lynden-Bell et al. 1988; DFB91: Dressler et al. 1991; B96: Bagglely 1996; JFK96: Jørgensen et al. 1996; L97: Lucey et al. 1997; D97: D’Onofrio et al. 1997; H97: Hudson et al. 1997; GFB01: Gibbons et al. 2001. 

Note b — (1) They used $\log D_\sigma = a \log \sigma + b$ with $D_\sigma$ in arcsec. Using $D_\sigma = \log (i_d R)$, where $i_d$ is in 0.1, one must add $\log R_{\text{CMB}} - 0.778$ to their zero point. (2) As in (1), but substitute $R_{\text{CMB}}$ for $R_{\phi}$, (3) The uncertainty in the distances determined using the inverse relation is $a \times \epsilon$.
the scatter of the composite distance relation and \( N \) is the number of galaxies observed in the cluster. The error in the cluster redshift is estimated as \( \sigma_{\text{cl}}/(N')^{1/2} \), where \( \sigma_{\text{cl}} \) is the velocity dispersion of the cluster and \( N' \) is the number of galaxies in the group catalog.

The distribution shown in the upper panel of Figure 9, which includes all 28 clusters, has an error-weighted mean of \( 151 \pm 75 \) km s\(^{-1}\), with a scatter of \( 399 \pm 73 \) km s\(^{-1}\). The lower panel shows that there are three obvious outliers: A2634, AS 639, and Cen 45, all based on data from the literature. Other clusters with large (>2 \( \sigma \)) peculiar velocities are Cen 30 (500 \( \pm 153 \) km s\(^{-1}\)), AS 714 (559 \( \pm 245 \) km s\(^{-1}\)), and AS 753 (812 \( \pm 204 \) km s\(^{-1}\)); all are located near the Great Attractor. If these clusters are removed from the sample, the mean peculiar velocity of the remaining 22 clusters is \( 71 \pm 51 \) and the rms one-dimensional cluster velocity is \( 239 \pm 46 \) km s\(^{-1}\). This is comparable to what is measured from the SCI sample (G97), \( 266 \pm 30 \) km s\(^{-1}\) (Giovanelli 1998). This small one-dimensional rms cluster velocity has important implications for cosmological parameters (e.g., Giovanelli 1998; Borgani et al. 2000b).

Notes to additional problematical clusters can be found in Appendix A. Most of these clusters appear to suffer from the effects of substructure; they have a history of discrepant peculiar motions reported in the literature.

The large overlap between our cluster sample and the literature allows a global comparison of the measured peculiar velocities. Figure 10 shows our cluster peculiar velocities \((v_p)\), computed using the direct \( D_n-\sigma \) relation, for the clusters we have in common with Jørgensen et al. (1996) (10 clusters), SCI (11 clusters), Hudson et al. (1997) (15 clusters), and Gibbons et al. (2001) (15 clusters). This figure shows that except for A2634, A194, and AS 753 our measurements of the cluster peculiar velocities are in good agreement with those reported in the literature when the measurement errors are taken into account. We find mean differences of \( 79 \pm 91 \) (Jørgensen et al. 1996), \( 182 \pm 94 \) (SCI), \( -9 \pm 96 \) (Hudson et al.
1997), and $-53 \pm 93$ km s$^{-1}$ (Gibbons et al. 2001). All clusters are in the same rest frame to within 2 $\sigma$. This agreement shows consistency between different determinations of cluster distances (e.g., those based on the $D_n$-$\sigma$ and/or FP relations) and, more importantly, consistency with the TF relation for spiral galaxies.

Figure 11 compares the peculiar velocities computed using the bivariate relation, corrected for selection bias (left panel), and the inverse relation (right panel) with those determined using the direct relation. The mean differences are $-43 \pm 32$ km s$^{-1}$ with a scatter of 74 km s$^{-1}$ and $-58 \pm 38$ km s$^{-1}$ with a scatter of 89 km s$^{-1}$ for the bivariate and inverse relations, respectively. This shows that the peculiar velocities of the clusters are largely insensitive to the fitting procedure, whether the direct, bivariate, or inverse relation is used.

5. DEPENDENCE OF THE DISTANCE RELATION ON GALAXY PROPERTIES

To use the composite $D_n$-$\sigma$ relation as a distance indicator, we should demonstrate that systematic cluster-to-cluster differences are small and that the computed cluster distances are unaffected by differences in the morphological mix of the galaxy population, different stellar populations, or other cluster properties. Furthermore, the measured peculiar velocities must be free of any other systematic effects such as extinction and contamination by interlopers. This can be tested by examining the residuals from the distance relation that, for our data, exceed the estimated measurement errors of the $D_n$-$\sigma$ parameters. (Note that testing to see if the distance relation depends on whether or not the galaxies are in clusters or in less dense environments is not...
the subject of this paper.) All the tests below suggest that cluster-to-cluster variations are indeed small.

5.1. Results for Individual Clusters

Figure 12 plots the measured values of $D_n$ and $\sigma$ for the galaxies in each cluster along with the (incompleteness-corrected) fit (dashed line) and the composite template relation (solid line). The parameters for the individual fits are given in Table 6: column (1) gives the cluster name; column (2) gives the number of galaxies entering the $D_n$-\$ relation; column (3) gives the slope and its error; column (4) gives the mean scatter in the data points relative to the individual fit; column (5) gives the zero-point offset between the individual and the composite template relation (eq. [9]); column (6) gives the scatter relative to the fit obtained using the slope of the composite template relation but allowing the individual zero point to vary; column (7) gives the intrinsic scatter computed using the scatter listed in column (6) and the errors of the measured parameters; and column (8) gives the fraction of the cluster galaxies in the observed sample that are elliptical galaxies [$F_E = N_E / (N_E + N_{S0})$].

Figure 12 shows that for most clusters the individual fits have nearly the same slope as the template. The figure also shows the benefit of combining all the data because the slope for the poorer systems in the sample is poorly determined. Significant departures ($\Delta \alpha > 0.2$) are seen for HMS 0122+3305, A1367, HG 50, A2199, Doradus, A3381, AS 639, Cen 30, and AS 714. The main cause for the tilt of an individual $D_n$-\$ relation is the small number of galaxies in the cluster. (For example, Table 6 shows that the tilt of the individual $D_n$-\$ relations does not correlate with the fraction of elliptical galaxies in the cluster, although some incorrect morphological classifications may still be present.) Indeed, even a single galaxy can cause a significant deviation from equation (9). As discussed above (see also Appendix A), many of these clusters show large motions. In addition, recall that A2199 and Cen 30 are parts of two-component systems (A2199/A2197 and Cen 30/Cen 45).

In general, Table 6 shows that the individual fit does not improve the scatter significantly and that variations are likely due to poor statistics. The source of the intrinsic scatter is still not understood, though the largest contributions to it probably arise from intrinsic differences in the dynamical structures of the cluster galaxies rather than from errors in the photometry and spectroscopy. Whether these intrinsic differences produce systematic errors in the distance determination is unknown. However, because we treat the thickness of the relation as an uncertainty in the derived distance, the impact of such scatter should not alter our conclusions about large-scale motions in the universe (see also § 5.2).

Figure 13 can be used to examine the impact of interlopers. It shows the residual of each galaxy from the distance relation as a function of the difference between the galaxy's redshift and that of the parent cluster. Field galaxies contaminating the sample would lie along the 45° line shown in each panel—no such effect is seen. This is a consequence of

| Name               | $n_{gal}$ | $l$  | $b$  | $\epsilon^{\text{CMB}}$ | $R$   | $\epsilon^{\text{CMB}}$ |
|--------------------|-----------|------|------|--------------------------|-------|--------------------------|
| 7S 21              | 7         | 113.784 | -40.018 | 5500 ± 90 | 5542 ± 441 | -41 ± 451 |
| Pisces             | 21        | 127.243 | -30.185 | 4715 ± 89  | 4626 ± 213 | 89 ± 231  |
| HMS 0122+3305     | 10        | 130.513 | -26.787 | 4600 ± 108 | 4525 ± 301 | 74 ± 320  |
| A262               | 8         | 136.999 | -29.049 | 4725 ± 93  | 4396 ± 328 | 328 ± 341 |
| A347               | 7         | 141.124 | -17.896 | 5301 ± 111 | 4808 ± 383 | 492 ± 398 |
| Perseus            | 26        | 150.382 | -13.382 | 4799 ± 133 | 4490 ± 185 | 309 ± 228 |
| A539               | 14        | 195.698 | -17.717 | 8636 ± 75  | 8119 ± 457 | 516 ± 464 |
| A1367              | 6         | 234.292 | 73.052  | 6807 ± 94  | 6989 ± 602 | -181 ± 609 |
| Virgo              | 44        | 283.871 | 74.200  | 1427 ± 49  | 1208 ± 38  | 219 ± 62  |
| Coma               | 80        | 58.301  | 88.285  | 7278 ± 75  | 7351 ± 173 | -72 ± 189 |
| HG 50              | 7         | 0.458   | 49.270  | 1905 ± 83  | 2240 ± 178 | -334 ± 197 |
| A2199              | 13        | 62.885  | 43.906  | 9108 ± 111 | 9069 ± 530 | 39 ± 542  |
| Pegasus            | 4         | 87.892  | -48.241 | 3202 ± 108 | 3635 ± 383 | -433 ± 398 |
| A2634              | 12        | 103.402 | -33.161 | 8975 ± 112 | 10762 ± 655 | -1787 ± 664 |
| A194               | 15        | 142.860 | -62.908 | 5074 ± 60  | 5079 ± 276 | -5 ± 283  |
| Fornax             | 18        | 236.241 | -54.096 | 1330 ± 36  | 1234 ± 61  | 96 ± 71   |
| Erindus            | 13        | 212.165 | -51.577 | 1488 ± 28  | 1827 ± 106 | -339 ± 110 |
| Doradus            | 4         | 260.209 | -47.227 | 1073 ± 36  | 1028 ± 108 | 45 ± 114  |
| A3381              | 6         | 240.293 | -22.697 | 11472 ± 65 | 10544 ± 908 | 927 ± 910  |
| Hydra              | 39        | 269.707 | 26.334  | 4055 ± 95  | 4103 ± 138 | -47 ± 168 |
| AS 639             | 6         | 280.534 | 10.908  | 6526 ± 93  | 4910 ± 423 | 1615 ± 433 |
| Cen 45             | 8         | 302.553 | 21.659  | 4931 ± 110 | 3013 ± 224 | 1918 ± 250 |
| Cen 30             | 21        | 302.023 | 21.852  | 3313 ± 82  | 2812 ± 129 | 500 ± 153 |
| AS 714             | 7         | 302.802 | 36.309  | 3576 ± 49  | 3017 ± 240 | 559 ± 245 |
| Klem 27            | 10        | 317.338 | 30.639  | 4881 ± 102 | 4402 ± 293 | 479 ± 310 |
| AS 793             | 18        | 319.160 | 26.744  | 4421 ± 97  | 3608 ± 179 | 812 ± 204 |
| Pavo II            | 12        | 352.191 | -23.755 | 4266 ± 63  | 4285 ± 261 | -18 ± 268 |
| Klem 44            | 18        | 25.336  | -75.807 | 8162 ± 88  | 8369 ± 416 | -206 ± 425 |

Note:—See Appendix A.
our membership assignment and the fact that early types are more likely to reside at the central regions of clusters.

To study the effects of morphology, we split the sample into elliptical ($T / C > 3$) and S0 galaxies ($T = 2$), using the Lauberts & Valentijn (1989) classifications. First, we computed the $D_{n} - C$ relation for the E and S0 galaxies separately. Figure 14 shows these relations for the elliptical (left panel) and S0 galaxies (right panel); the difference in the slope of the two distance relations is $0.071 \pm 0.059$, which is significant at the less than 2 $\sigma$ level, and the scatter is comparable. Second, we determined the relative shifts that were required if a linear relation of the same slope as of the composite template relation ($a = 1.203$; see eq. [9]) was to fit the relation in each of the subsamples. The difference in the intercept is not statistically significant, and the scatter is comparable. These results justify our neglect of any morphological biases ($\S$ 3.1).

To test further the above result, one could instead consider the residuals in the $D_{n} - C$ relation as a function of the $D / B$ ratio. Unfortunately, in practice, $D / B$ is available only for those galaxies in our sample that were observed by us; the data compiled from the literature used one-component models to derive global photometric parameters. The 223 cluster galaxies for which we have our own photometric measurements show no correlation between the residuals of the $D_{n} - C$ relation and the $D / B$ ratio (see Fig. 15).

5.2. Stellar Populations

Earlier in this paper we found that the scatter of galaxies relative to the template distance relation is roughly twice...
what can be accounted for by measurement errors. The additional scatter has been attributed, by several authors, to differences in stellar populations. In the context of distance measurements, we must check if these differences can lead to systematic errors in the distance and therefore to spurious peculiar velocities.

To study the effects of different stellar populations, we use the \( \text{Mg} \, 2-\sigma \) relation, which is supposed to be distance independent. It was computed for all galaxies in the sample after sorting them according to their morphological types (see Fig. 16). The parameters of the orthogonal fits to the \( \text{Mg} \, 2-\sigma \) relation are given in Table 7: column (1) gives the morphological types that are in the sample; column (2) gives the number of galaxies; column (3) gives the slope computed from the whole sample and its error; column (4) gives the zero point and its error computed by fixing the slope to the value reported in column (3); and column (5) gives the scatter relative to the relation. Note that the coefficients describing the linear fit obtained here differ slightly (<1 \( \sigma \)) from those of Bernardi et al. (1998). This is because we now include \( \text{Mg} \, 2 \) measurements from other authors, scaling them to our system. The table shows small differences in the zero point between elliptical and S0 galaxies, although the

![Fig. 9.](image)

**Fig. 9.**—Distribution of cluster peculiar velocities (upper panel) and cluster peculiar velocities vs. the estimated distances (lower panel). Filled circles represent the “distant” clusters used for the final calibration of the \( D_\alpha-\sigma \) relation, open circles represent either nearby clusters or clusters that have suspiciously large peculiar velocities, and triangles indicate clusters that were not observed by our survey.

**Table 6**

| Cluster       | \( N_{\text{gal}} \) | \( a \)  | \( \epsilon \) | \( \Delta \beta \) | \( \epsilon_\alpha \) | \( \epsilon_{\text{intr}} \) | \( F_\alpha \) |
|---------------|----------------------|---------|----------------|-----------------|----------------|----------------|-----------|
| 7S 21         | 7                    | 1.150 ± 0.093 | 0.089 | 0.001 | 0.090 | 0.074 | 0.43         |
| Piscis        | 21                   | 1.200 ± 0.035 | 0.069 | -0.006 | 0.068 | 0.053 | 0.57         |
| HMS 0122 + 3305 | 10                | 1.523 ± 0.080 | 0.065 | -0.028 | 0.074 | 0.062 | 0.50         |
| A262          | 8                    | 1.251 ± 0.085 | 0.089 | -0.013 | 0.090 | 0.078 | 0.88         |
| A347          | 7                    | 1.246 ± 0.206 | 0.046 | -0.013 | 0.045 | 0.012 | 0.57         |
| Perseus       | 26                   | 1.220 ± 0.047 | 0.082 | -0.027 | 0.080 | 0.068 | 0.62         |
| A539          | 14                   | 1.053 ± 0.062 | 0.071 | -0.007 | 0.072 | 0.060 | 0.14         |
| A1367         | 6                    | 1.883 ± 0.295 | 0.053 | -0.034 | 0.063 | 0.044 | 1.00         |
| Virgo         | 44                   | 1.168 ± 0.011 | 0.115 | 0.000 | 0.115 | 0.108 | 0.64         |
| Coma          | 80                   | 1.271 ± 0.012 | 0.071 | -0.002 | 0.071 | 0.063 | 0.61         |
| HG 50         | 7                    | 1.002 ± 0.127 | 0.091 | -0.018 | 0.093 | 0.080 | 0.71         |
| A2199         | 13                   | 1.549 ± 0.069 | 0.125 | 0.006 | 0.131 | 0.125 | 0.85         |
| Pegasus       | 4                    | 1.315 ± 0.251 | 0.057 | -0.027 | 0.058 | 0.013 | 0.50         |
| A2634         | 12                   | 1.323 ± 0.112 | 0.058 | -0.014 | 0.058 | 0.046 | 0.42         |
| A194          | 15                   | 1.120 ± 0.038 | 0.079 | 0.012 | 0.080 | 0.070 | 0.47         |
| Fornax        | 18                   | 1.153 ± 0.022 | 0.108 | 0.017 | 0.108 | 0.101 | 0.50         |
| Eridanus      | 13                   | 1.038 ± 0.029 | 0.093 | 0.035 | 0.101 | 0.093 | 0.54         |
| Doradus       | 4                    | 0.793 ± 0.179 | 0.042 | 0.001 | 0.066 | 0.037 | 0.50         |
| A3381         | 6                    | 1.408 ± 0.242 | 0.065 | -0.016 | 0.067 | 0.045 | 0.33         |
| Hydra         | 39                   | 1.183 ± 0.017 | 0.108 | 0.009 | 0.108 | 0.101 | 0.28         |
| AS 639        | 6                    | 0.927 ± 0.118 | 0.055 | -0.009 | 0.068 | 0.047 | 0.50         |
| Cen 45        | 8                    | 1.140 ± 0.086 | 0.047 | 0.001 | 0.047 | 0.022 | 0.12         |
| Cen 30        | 21                   | 1.623 ± 0.031 | 0.111 | -0.019 | 0.130 | 0.125 | 0.52         |
| AS 714        | 7                    | 0.956 ± 0.166 | 0.048 | -0.001 | 0.051 | 0.023 | 0.00         |
| Klem 27       | 10                   | 1.252 ± 0.060 | 0.088 | -0.009 | 0.089 | 0.080 | 0.50         |
| AS 753        | 18                   | 1.090 ± 0.031 | 0.097 | -0.006 | 0.100 | 0.092 | 0.17         |
| Pavo II       | 12                   | 1.151 ± 0.061 | 0.086 | -0.006 | 0.086 | 0.079 | 0.67         |
| Klem 44       | 18                   | 1.105 ± 0.030 | 0.109 | 0.008 | 0.111 | 0.103 | 0.44         |

**Table 7**

| Sample | \( N_{\text{gal}} \) | \( a \)  | \( b \)  | \( \epsilon \) |
|--------|----------------------|---------|---------|------------|
| All........ | 369 | 0.226 ± 0.014 | -0.227 ± 0.010 | 0.021 |
| E ........... | 186 | -0.226 ± 0.014 | 0.019 |
| S0 .......... | 183 | -0.230 ± 0.015 | 0.023 |
Fig. 10.—Cluster peculiar velocities obtained using eq. (9) vs. the values computed by Jørgensen et al. (1996) (JFK96), G97 (SCI), Hudson et al. (1997) (H), and Gibbons et al. (2001) (G).

Fig. 11.—Cluster peculiar velocities obtained using the direct $D_L$-$\sigma$ relation (eq. [9]) vs. the values computed using the bivariate (left panel) and the inverse relations (right panel).
S0 galaxies have a larger scatter than elliptical galaxies. This is partially due to the small number of galaxies with low velocity dispersions in both subsamples. Although one expects S0 galaxies to form a less uniform class of objects than elliptical galaxies, the differences we find are small. Therefore, this analysis suggests that the $D_n/C_27$ relation does not depend on differences in stellar populations.

If differences in stellar populations were important in estimating distances, one would expect correlations between the $D_n/C_27$ and $\Delta M_{g2}/C_27$ residuals, $\Delta(D_n/C_27)$ and $\Delta M_{g2}$, respectively, since the latter should reflect either age or metallicity differences. The left panels of Figure 17 show $\Delta(D_n/C_27)$ versus $\Delta M_{g2}$ for the cluster galaxies as a whole (upper panel), for the elliptical galaxies (middle panel), and for S0 galaxies (lower panel), while the right panels show $\Delta(D_n/C_27)$ as a function of the $M_{g2}$ line index. As can be seen, there is no obvious correlation between these parameters. Applying the Spearman rank test to the data shown in the various panels, we find that the rank-order correlation coefficients vary from 0.10 to $-0.15$, implying significance levels greater than 0.8, thereby confirming the lack of any significant correlation between these quantities. These results are in agreement with the conclusions of previous studies (e.g., Jørgensen et al. 1996; Colless et al. 1999).

Figure 18 shows the data points for each cluster and the composite $M_{g2}/C_27$ relation (solid line) given by the sample as a whole. Open circles indicate S0 galaxies, and filled circles indicate elliptical galaxies. From these panels, it is evident that S0 galaxies depart more from the composite relation than elliptical galaxies, especially at smaller velocity dispersions. Nevertheless, most galaxies do lie along the globally derived relation. There are some

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**Fig. 12.—** Individual cluster $D_n/C_27$ relations obtained by fitting the bias-corrected data points of the cluster (dashed line). The solid line in all panels shows the template distance relation given by eq. (9).
exceptions that were observed by other authors; these are listed in Appendix B.

Our results suggest that differences in stellar populations do not influence the distance relation enough to mimic peculiar motions. Furthermore, none of the most discrepant peculiar velocities discussed in the previous section show evidence that their velocities are caused by stellar population effects.

5.3. Environment

In the literature, there is concern that there may be cluster-to-cluster environmental differences in the \( D_n - \sigma \) method, and here we examine this possibility. Figure 19 shows the distribution of the intrinsic scatter \( \epsilon_{\text{intr}} \) and the slope of the fit as a function of (1) the measured velocity dispersion of the cluster \( \sigma_{cl} \) and (2) the logarithm of the ratio \( \sigma_{cl}^2/R_p \), where \( R_p \) is the pair radius defined by Ramella, Geller, & Huchra (1989) and \( \sigma_{cl}^2/R_p \) is a rough measure of the projected cluster surface density. Clusters/groups with poorly defined slopes (those with an error in the slope \( \geq 0.1 \); see Table 6 and Fig. 12) are shown as crosses. Also represented by crosses are the points corresponding to the systems HMS 0122+3305, A2199, and Cen 30, where possible membership assignment problems may affect the determination of the slope (see B02). Seven clusters from the literature, for which values for \( R_p \) are not available, have not been included in the right panels of the figure (see B02). Applying the Spearman rank test to the whole sample confirms that there is no obvious correlation between the intrinsic scatter and the parameters that characterize the global properties of the clusters. The rank-order correlation
coefficients are 0.01 and −0.05 with significance levels of 0.96 and 0.77, respectively. On the other hand, a similar analysis of the data shown in the lower panels of Figure 19 might lead one to suspect that the slope of the $D_n$-$\sigma$ relation depends both on the velocity dispersion and on the surface density. Taking all the available data points into consideration, the Spearman rank test gives a rank-order correlation coefficient of 0.60 with significance levels of 0.005 and 0.010 in both cases, indicating a strong correlation between the slope and $\sigma_{cl}$ or $\log \sigma_{cl}^2/\langle R_p \rangle$. However, the slope of individual cluster/group relations is in many cases poorly determined either because of the small number of measured cluster members or because of interlopers. In fact, if systems with large errors in the slope (nine systems) and cases in which the slope could be affected by the presence of interlopers (three systems) are discarded, the Spearman rank-order correlation coefficient decreases to 0.40, corresponding to a significance level of 0.15 for both relations; this shows that the correlation between the slope of the $D_n$-$\sigma$ relation and the velocity dispersion or central surface density is not significant. Clearly, a more definite test of this hypothesis requires considerable more data per cluster than currently available. We should also point out that the Spearman rank test also shows that there is no obvious correlation between either the intrinsic scatter or the slope of the cluster’s individual $D_n$-$\sigma$ relation and the number of observed galaxies. The derived rank-order correlation coefficients are 0.32 and 0.11, yielding significance levels of 0.4 and 0.8, respectively. It should be emphasized at this point that the good agreement in the peculiar velocity field obtained from spirals and elliptical galaxies gives further support to the hypothesis of a universal distance relation.

Gibbons et al. (2001) used 20 clusters, of which 15 are in common with us, to argue that the amplitude of the measured peculiar velocity correlates with the scatter of the distance relation. The left panel in Figure 20 shows the cluster peculiar velocities of all the clusters in our sample as a function of the amplitude, $\epsilon_{\text{intr}}$, of the intrinsic scatter of the individual $D_n$-$\sigma$ relations. The right panel shows the fraction of elliptical to early-type galaxies $[N_E/(N_E + N_{S0})]$ versus $\epsilon_{\text{intr}}$. These plots are similar to those shown by Gibbons et al. (2001). However, in contrast, we find no significant correlations in either relation. The Spearman rank test gives a rank-order correlation coefficient of 0.21 with a significance level of 0.48, yielding significance levels of 0.4 and 0.8, respectively. It should be emphasized at this point that the good agreement in the peculiar velocity field obtained from spirals and elliptical galaxies gives further support to the hypothesis of a universal distance relation.

We conclude that cases of poor fits are more likely to be due to observational limitations rather than reflecting intrinsically different physical properties. In summary, we find no compelling evidence that the peculiar velocities are spurious artifacts. Rather, we believe that our quoted velocities do measure the motion of clusters relative to the Hubble flow.
6. SUMMARY

Using new and previously published data for 452 galaxies in 28 clusters, we have derived a bias-corrected \( D_n - \sigma \) relation. It can be used to measure relative distances of galaxies in the recently completed survey of early-type galaxies (da Costa et al. 2000b) and to map the peculiar velocity field. Our main conclusions are as follows:

1. The slope obtained by combining data for all clusters/groups does not differ significantly from previous determinations.

2. The scatter is found to be \( \sim 0.085 \) dex, implying a distance error of about 20% per galaxy, comparable to the error of FP relations. Note that \( D_n \) is, in general, less sensitive to seeing and easier to compute (e.g., fits to light profiles are not required).

3. Our cluster peculiar velocities are in good agreement with other determinations, in particular, with those based on spiral TF distances, further supporting the validity of the distance indicators.

4. As in previous work, we find no evidence for systematic effects playing a role in the computed peculiar velocities. We believe that the peculiar velocities we present here are not artifacts but rather a true measure of the clusters’ motions relative to the Hubble flow.

5. Of the 28 clusters in the sample, six show suspiciously large peculiar velocities (both infall and outflow). Five of these are likely due to small-scale dynamical effects or contamination by other components. The remaining one is at low Galactic latitude and may suffer from absorption effects. Eliminating these clusters, we find that the cluster one-dimensional rms velocity is relatively small, \( 239 \pm 46 \) km s\(^{-1}\), suggesting a fairly quiescent velocity field, consistent with the estimate obtained from the TF data.

The distance relations derived here have been used in previous papers of this series (da Costa et al. 2000a; Borgani et al. 2000a; Nusser et al. 2001; Zaroubi et al. 2001) to analyze the peculiar velocity field traced by early-type galaxies. This sample of early types, comparable in size to the SFI sample of field spirals (Haynes et al. 1999a, 1999b), allows an independent analysis of the characteristics of the local velocity field because it uses a different distance relation and test particles that probe a different set of density regimes. The good agreement between our early-type cluster sample and the SCI spiral sample suggests that it should be possible to merge the ENEAR and SFI redshift surveys. This will provide the largest and most homogeneous all-sky sample of nearby galaxies available for cosmic flow studies and will allow the universality of the results presented here to be checked directly.

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APPENDIX A

NOTES ON THE MOST PECULIAR CLUSTER VELOCITIES

A2634.—This is one of the most distant clusters in the ENEARc sample ($cz \sim 9000$ km s$^{-1}$) and has a large peculiar velocity (it is more than 2 $\sigma$ from the mean defined by our full sample). We have no measurements of our own for this cluster. Lucey et al. (1997) reobserved some of its galaxies and concluded that the original values of the central velocity dispersion were underestimated. This partially accounts for its large infall velocity. Here the peculiar velocity was determined using the Lucey et al. (1997) estimates.

Fig. 17.—*Left panels:* Residuals with respect to the mean $D_n$-$\sigma$ relation vs. residuals with respect to the mean Mg$_{2}$-$\sigma$ relation for the cluster sample as a whole (*upper panel*), for elliptical galaxies (*middle panel*), and for S0 galaxies (*lower panel*). *Right panels:* Same as in the left panels, but now for the residuals of the $D_n$-$\sigma$ relation vs. the measured Mg$_{2}$ index.
converted to our system. Although our value (approximately $-1787 \, \text{km s}^{-1}$) is smaller than the number reported in Faber et al. (1989) and Lucey et al. (1991) ($-3400 \pm 600 \, \text{km s}^{-1}$), it also disagrees with the more recent estimate from Lucey et al. (1997) and the estimate from the SCI sample of cluster spirals (G97). The Lucey et al. (1997) value is significantly smaller ($\lesssim -700 \, \text{km s}^{-1}$) than ours, with the actual value depending on the distance relation used (FP or $D_n$-$\sigma$). In the SCI sample this cluster is nearly at rest relative to the Hubble flow. Hudson et al. (1997), using a subsample of the Lucey et al. (1997) data set and the FP relation in their paper, also find a small infall. However, using the slope of their $D_n$-$\sigma$ relation and a zero point derived from the peculiar velocities of the other two clusters we have in common with them, A347 and 7S 21 (because the zero point of the $D_n$-$\sigma$ relation is not reported in that paper), yields a peculiar velocity of $-1582 \pm 635 \, \text{km s}^{-1}$ for A2634, comparable to our value. We also note that A2634 has a nearby companion (A2166) at approximately the same redshift, which may affect membership assignment and may explain the large variations in its measured peculiar velocity.

A3381.—This is the most distant cluster/group in the ENEARc sample ($cz_{\text{CMB}} = 11472 \pm 65 \, \text{km s}^{-1}$) with only six early-type galaxies. This cluster was originally studied by Jørgensen et al. (1996), who reported a peculiar velocity of $667 \pm 698 \, \text{km s}^{-1}$.

A3639.—At low Galactic latitude ($b \sim 10^\circ$), A3639 was originally studied by Jørgensen et al. (1996), who found it outflowing at $1295 \pm 359 \, \text{km s}^{-1}$. Correcting to our standard system, we find an amplitude of $1615 \pm 433 \, \text{km s}^{-1}$. This may reflect differences in the galaxy sample, since we have removed ESO 264-G024, 1037-4605, ESO 264-IG030

Fig. 18.—$M_g$ index of each cluster member galaxy vs. its velocity dispersion. Open circles indicate S0 galaxies, while filled circles indicate elliptical galaxies. The solid line shows the derived composite $M_g$-$\sigma$ relation.
NED 03, and ESO 264-IG 030 NED 02 from it (see Table 8 in B02). It may also reflect differences in the adopted distance relations. Jørgensen et al. (1996) argued that this large amplitude was partially due to stellar population differences \((\leq 5.2)\). Using the correlation between the Mg 2 line index and the central velocity dispersion, they argued that the amplitude of the motion was smaller than \(\sqrt{C_24} / \sqrt{C_6} 879 \text{ km s}^{-1} / \sqrt{C_24} \) Recently, Jørgensen & Jønch-Sørensen (1998), using additional data, found a peculiar velocity of \(\sqrt{C_6} 350 \text{ km s}^{-1} / \sqrt{C_24} \). They argued that this is also an overestimate because of evidence for an apparently younger stellar population. This cluster lies so close to the Galactic plane that uncertainties in absorption correction may be large; these may lead to artificially high values of the peculiar velocity.

Cen 30 and Cen 45.—Their large peculiar velocities can be partially explained by the fact that they lie along the same line of sight and are part of a complex structure. In Figure 7, Cen 30 shows a bimodal distance distribution because it is difficult to assign galaxies to the different clumps. While clearly seen in the distance distribution, the bimodality is not evident in the redshift distribution in Figure 8. The large positive peculiar velocity of Cen 45 is likely caused by its

\[
\sigma_{cl} \text{ [kms}^{-1}] \quad \frac{2 \log \sigma_{cl} - \log R_p}{2}
\]

Fig. 19.—rms scatter (upper panels) and the slope (bottom panels) of the individual cluster \(D_n-\sigma \) relations vs. the measured velocity dispersion of the cluster \(\sigma_{cl} \) (left panels) and the logarithm of the ratio \(\sigma_{cl} / R_p \) (right panels), where \(R_p \) is the pair radius defined by Ramella et al. (1989). Seven clusters taken from the literature, which were not identified by the finding algorithm, are not included in the right panels. Crosses represent clusters/groups with a large error in the slope \((\leq 0.1); \text{ see Table 6}) \) or that exhibit clear evidence of either spatial substructure or distinct galaxy populations (HMS 0122+3305, A2199, Cen 30), and filled circles represent clusters/groups with reliable \(D_n-\sigma \) fits.

\[
\begin{align*}
F_E &= N_E / (N_E + N_S) \\
N_E &= \text{number of elliptical galaxies} \\
N_S &= \text{number of spiral galaxies}
\end{align*}
\]

Fig. 20.—Left panel: Peculiar velocities of the 28 clusters as a function of the amplitude of the scatter of the individual \(D_n-\sigma \) relations of each cluster. Right panel: Fraction of elliptical to early-type galaxies \(F_E = N_E / (N_E + N_S) \), for each cluster, vs. the scatter of the individual \(D_n-\sigma \) relations.
infall toward the more massive component of the system (e.g., Lucey & Carter 1988). Given the complexity of the Centaurus system, one should be cautious when using Cen 30 and Cen 45.

**AS 714.**—This cluster, with suspiciously large amplitude, has 19 members, close to the minimum number required to be included in the cluster sample. We targeted all eight early types in it, of which six are lenticulars. One of these was excluded from the cluster sample used to derive the $D_{\text{v}}$ relationship because it appears to be spiral (see B02, Table 8). The measured peculiar velocity, $559 \pm 245$ km s$^{-1}$, is high. However, the group is located in the direction of the Great Attractor (GA), which may account for the large amplitude (as for Cen 30). Because of the complexity of the region, the large peculiar velocity can also arise from small-scale dynamical effects, such as those in Cen 45.

**AS 753.**—This cluster, also in the GA region, shows a large positive peculiar velocity of $812 \pm 204$, which is significantly larger than the $279 \pm 182$ km s$^{-1}$ obtained by Jørgensen et al. (1996). The difference between the Jørgensen et al. (1996) value and ours is partially due to the choices of the galaxy sample, the distance relation, and our weighting procedures, illustrating how systematic rather than random errors can sometimes be responsible for significant differences in the measured peculiar velocity of individual clusters.

### APPENDIX B

NOTES ON THE Mg$_{2}$-$\sigma$ RELATION

The following galaxies lie off the Mg$_{2}$-$\sigma$ relation shown in Figure 18.

**Perseus.**—PGC 012423 (Smith et al. 1997) shows a higher Mg$_{2}$ index than that expected from the Mg$_{2}$-$\sigma$ relation.

**A339.**—Three galaxies (CGCG 421-015, CGCG 421-017, and 0514-0619a from Jørgensen et al. 1995b) show a lower Mg$_{2}$ index. They are faint galaxies in a crowded background.

**A3381.**—PGC 018554 (Jørgensen et al. 1995b) has a lower Mg$_{2}$ index than expected. The spectrum of this galaxy may be affected by the light of a nearby bright star.

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