IR spectral fingerprint of carbon monoxide in interstellar water–ice models

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Abstract
Carbon monoxide (CO) is the second most abundant molecule in the gas phase of the interstellar medium. In dense molecular clouds, it is also present in the solid phase as a constituent of the mixed water-dominated ices covering dust grains. Its presence in the solid phase is inferred from its infrared (IR) signals. In experimental observations of solid CO/water mixed samples, its IR frequency splits into two components, giving rise to a blue- and redshifted band. However, in astronomical observations, the former has never been observed. Several attempts have been carried out to explain this peculiar behaviour, but the question still remains open. In this work, we resorted to pure quantum mechanical simulations in order to shed some light on this problem. We adopted different periodic models simulating the CO/H$_2$O ice system, such as single and multiple CO adsorption on water–ice surfaces, CO entrapped into water cages, and pure CO:H$_2$O mixed ices. We also simulated pure solid CO. The detailed analysis of our data revealed how the quadrupolar character of CO and the dispersive forces with water ice determine the energetic of the CO/H$_2$O interaction, as well as the CO spectroscopic behaviour. Our data suggest that the blueshifted peak can be assigned to CO interacting via the C atom with dangling H atoms of the water ice, while the redshifted one can actually be the result of CO involved in different reciprocal interactions with the water matrix. We also provide a possible explanation for the lack of the blueshifted peak in astronomical spectra. Our aim is not to provide a full account of the various interstellar ices, but rather to elucidate the sensitivity of the CO spectral features to different water–ice environments.

Key words: astrochemistry – molecular processes – methods: numerical – techniques: spectroscopic – ISM: lines and bands – ISM: molecules.

1 INTRODUCTION
In cold (≲ 100K) molecular clouds, carbon monoxide (CO) is the second most abundant gaseous molecule (after hydrogen, H$_2$). In addition, given its spectroscopic and chemical properties, CO is widely used as a proxy of H$_2$ of molecular gas in both galactic and extragalactic sources. However, in very cold (≲ 25K) and very dense (≥ 10$^4$ cm$^{-3}$) gas, CO molecules freeze out onto the interstellar grain surfaces, in the so-called grain mantles, and disappear from the gas (Caselli et al. 1999). The effect can be extreme in particularly dense regions, as prestellar cores and protoplanetary discs, where more than 90 per cent of CO is likely in the solid rather than gas phase (Bacmann et al. 2002; Favre et al. 2013). Therefore, solid CO may be the major reservoir of elemental carbon (not locked into carbonaceous grains) in a large fraction of the cold interstellar medium (ISM).

The way how CO freezes, namely in what structure (whether trapped in a water-rich matrix or in pure CO layers), affects the temperature at which it returns into the gas phase. In turn, this has important consequences on a variety of situations; for example, on the molecular deuteration of water and trace species (Ceccarelli et al. 2014), on the formation of methanol and more complex organic molecules in the protostellar phase (Garrod & Herbst 2006), or on
the composition of the gaseous giant planets (Oberg & Bergin 2016; Madhusudhan et al. 2017).

Therefore, understanding the molecular structure of the interstellar solid CO is of great importance. This can only be obtained by observations of the 4.647–4.682 µm (2152–2136 cm\(^{-1}\))\(^{12}\)CO and 4.780 µm (2092 cm\(^{-1}\))\(^{13}\)CO absorption features towards astronomical sources and their comparison with ”expected” solid CO spectra, based on laboratory and/or theoretical works.

From a theoretical point of view, grain mantles grow in molecular clouds and in the denser fragments of them, called prestellar cores. Water ice forms via hydrogenation of oxygen atoms that land on the grain surfaces (Dulieu et al. 2010; Oba et al. 2012) whereas CO ice results directly from the sticking of gaseous CO onto the grains. Observations and models clearly indicate that water ice is the most abundant component of the grain mantles and that it is formed before the CO ice (Taquet et al. 2013; Boogert, Gerakines & Whittet 2015). Therefore, in principle, CO molecules can be trapped into a water-rich matrix or in pure-CO layers. In addition, UV photons and cosmic rays are believed to affect the iced mantle structure (Oberg, van Dishoeck & Linnartz 2009). As a result of these several different phenomena, ice seems to be present in two rather different phases: (i) as a highly porous amorphous film containing several apolar molecules and characterized by rather large cavities (Al-Halabi et al. 2004b) and (ii) as a layered and more ordered highly hydrogenated phase distinguished by a significant concentration of dangling groups, i.e. water molecules at a surface or discontinuity that cannot complete their tetrahedrally hydrogen bonded network with their neighbouring molecules (Collings et al. 2003a).

The CO/H\(_2\)O interplay takes place in several different and peculiar environments, including low pressure and temperature, and is expected to include several competitive aspects such as surface–surface interactions, CO adsorption, migration into ice pores (Devlin 1992), and CO trapping into the ice matrices (Sandford et al. 1988; Schmitt, Greenberg & Grim 1989). Numerous experiments have addressed different aspects of the pure-H\(_2\)O and of the CO-containing ices, yielding a great number of data that prompt multiple and sometimes controversial issues. In particular, infrared (IR) spectral analysis is the main experimental instrument to investigate the physical chemistry of CO interacting with H\(_2\)O ice under interstellar-like conditions. However, it is sometimes difficult to orient in the plethora of experimental peaks, often faint and concentrated in a small portion of the spectrum. In this regard, quantum mechanical simulations can represent fundamental tools for deriving important structural and atomistic information. The present manuscript is intended to calculate, within the same theoretical first principle approach, the IR fingerprints of several different atomistic models for the CO/H\(_2\)O ice interaction in order to facilitate the comparison among the various spectroscopic signals and to provide tentative molecular models for each of them in a coherent and robust framework.

The paper is organized as follows. In Section 2, we briefly review the situation about the laboratory and theoretical works on the IR absorption bands observed in the astronomical sources and the open questions about the interpretation of these observations. In Section 3 we resume the adopted computational methods for static and dynamic calculations. In Section 4, we present and discuss the results concerning the different CO/H\(_2\)O ice systems we modelled in the light of the most recent experimental findings and theoretical analysis. General hypotheses on some controversial issues are suggested at each subsection. Finally, in Section 5, all results are briefly summarized and some perspectives are also drawn.

2 BRIEF REVIEW OF PREVIOUS STUDIES

In IR observations, gas-phase CO is detected at 2143 cm\(^{-1}\) (4.666 µm) due to its C–O stretching mode (Mina-Camilde, Manzanares & Caballero 1996). Therefore, the analysis of the ISM spectra focuses on the splitting and shifts of this signal towards lower (red) or higher (blue) wavenumbers. Thus, before any investigation, an overview of the detected IR signals for CO/H\(_2\)O systems along with their interpretations is done here.

As already mentioned in the Introduction section, two rather different ice structures are supposed to be present in the ISM, referred to as the apolar mixed (Ap) and the layered polar (Pr) phase. In order to reproduce both of them, as well as their interaction with carbon monoxide, different experimental settings and computational simulations have been performed over the years.

CO co-deposited in water matrices at low temperatures, as a model of Ap, presents a very peculiar band profile consisting of two main broad features: (i) an intense peak at 2138 cm\(^{-1}\) (4.677 µm), redshifted with respect to the gas phase and indicating a decreasing of the bond strength (and a corresponding increasing of the C–O bond length), and (ii) a secondary, less-intense, blueshifted peak at 2152 cm\(^{-1}\) (4.647 µm) corresponding to a shortening/increasing of the bond length/strength. This two-peak profile has been noted and discussed by many authors.

According to Sandford et al. (1988) and Jenniskens et al. (1995), the 2138 cm\(^{-1}\) band is due to substitutional CO molecules, i.e. CO replacing H\(_2\)O lattice molecules, while the 2152 cm\(^{-1}\) band is caused by interstitial CO, i.e. CO inside H\(_2\)O cages. An alternative interpretation attributes the 2152 cm\(^{-1}\) peak to CO molecules interacting with the OH dangling groups (Devlin 1992; Collings et al. 2003a), hereafter referred to as OH sites, while the 2138 cm\(^{-1}\) is assigned to CO diffused into micropores (Palumbo 1997).

Interstellar H\(_2\)O ice is believed to be predominantly amorphous, although its specific morphology remains an open question when referring to its porosity. To investigate on the possibility of CO migrating into ice lattice, Al-Halabi and coworkers (Al-Halabi et al. 2004b) have studied the interaction of CO with amorphous ice by means of classical trajectory calculations and have concluded that CO neither penetrates into the ice slab matrix nor diffuses into the surface valley. The eventuality of CO diffusion appears as a temperature effect (Palumbo 1997): when ice is warmed up to 30 K, CO mainly desorbs but few molecules are trapped in the H\(_2\)O porous structure giving rise to a narrowing of the band at 2138 cm\(^{-1}\) related to CO in water matrix. At temperatures higher than 80 K, the disappearance of the peak at 2152 cm\(^{-1}\) indicates the break of the OH–CO interaction. Interestingly, the annealing process of CO-rich ices produces a family of new adsorption features, at 2143 cm\(^{-1}\), that have been tentatively associated with the formation of clathrate-like structures (Sandford et al. 1988).

The CO-layered ice binary system Pr, although prepared following different experimental procedures and theoretical models, presents almost the same spectroscopic pattern. At low CO coverage, it can be seen as a gas–surface interaction. When CO concentration increases, a CO/H\(_2\)O interface appears and a layered CO structure starts to grow. To characterize this system, Allouche, Verlaque & Pourcin (1998) have performed a combined experimental/computational study of CO adsorption on ice surfaces. They measured an adsorption enthalpy of about 10 kJ mol\(^{-1}\) and unambiguously assigned the peak at 2152 cm\(^{-1}\) to CO interacting with OH. In the context of this tendency toward a long-range crystalline order, the peak at 2143 cm\(^{-1}\) has been assigned to the adsorption of linearly polarized light in the CO ice phase, which starts to
crystallize as the temperature increases (Pontoppidan et al. 2003). Finally, two combined effects are supposed to give rise to the 2139–2136 cm$^{-1}$ features: CO/CO interactions inside the rising CO multilayer as well as interactions involving dangling O atoms (O0) of the exposed water molecules occurring at the CO/H$_2$O interface (Manca, Roubin & Martin 2000; Martin, Manca & Roubin 2000). This band narrows as the temperature increases, due to the progressive CO desorption (Al-Halabi, van Dishoeck & Kroeis 2004a).

While nowadays there is a currently general agreement as regard the peak at 2152 cm$^{-1}$, present in both the apolar mixture and the hydrogenated ice surfaces, and assigned to CO adsorbed on the exposed water molecules occurring at the CO/H$_2$O interface (Manca, Roubin & Martin 2000; Martin, Manca & Roubin 2000). This band narrows as the temperature increases, due to the progressive CO desorption (Al-Halabi, van Dishoeck & Kroeis 2004a).

When Gaussian basis sets are used, as in CRYSTAL17, the basis set superposition error (BSSE) arises in the evaluation of the interaction energy (Boys & Bernardi 1970). One of the most common a posteriori correction is represented by the ‘Counterpoise’ (CP) method, which exploits ‘ghost functions’ (Davidson & Feller 1986). For all tested cases, we computed the BSSE-corrected interaction energy per adsorbed CO molecule ($\Delta E^\infty$) of the CO-ice complexes according to Equation (1)

$$\Delta E^\infty = \frac{\Delta E^\infty + \delta E + E_L - \text{BSSE}}{n}$$  \hspace{1cm} (1)

where $\Delta E^\infty$ is the deformation-free interaction energy, $\delta E = \delta E_1 + \delta E_2$ is the total deformation energy of both the ice model and the CO molecules, $E_L$ is the lateral interaction among different replicas of the periodic models, BSSE is the total BSSE contribution, and $n$ is the number of CO molecules in the complex. To fully account for quantum mechanical effects, we also computed the interaction enthalpies $\Delta H$ at 0 K, according to Equation (2)

$$\Delta H = \Delta E^\infty + AZPE$$  \hspace{1cm} (2)

where $AZPE$ is the difference (again per adsorbed CO molecule) between the zero-point energy (ZPE) of the $n$ CO molecules in the complexes (considered as a fragment of the whole structure) and $n$ times the ZPE for the isolated CO, unless otherwise specified. For a more detailed discussion of Equation (1) and of the assumptions underlying Equation (2), please see the ‘Computational details’ section of Appendix A.

We scaled all the harmonic computed CO stretching wavenumbers with a proper scaling factor defined as the ratio between the experimental (2143 cm$^{-1}$; Mina-Camilde et al. 1996) and our calculated (2118 cm$^{-1}$) value for the isolated CO molecule. This allows to partially recover for the intrinsic errors associated with the computational level we adopted and also to account for anharmonic effects.

The IR simulated spectra are defined as linear superpositions of Lorentzian functions centred at the different computed wavenumbers and weighted for the corresponding IR intensities. We choose the two arbitrary values of 3 and 15 cm$^{-1}$ as full width at half-maximum (FWHM) to account for the natural width of the experimental and astronomical measures.

When dealing with 2D periodic slab models, we evaluated the relative stability of a given slab model compared to the corresponding bulk through the calculation of the surface energy $E_S$ according to Equation (3)

$$E_S = \frac{E_{\text{slab}} - mE_{\text{bulk}}}{2A}$$  \hspace{1cm} (3)

where $E_{\text{slab}}$ and $E_{\text{bulk}}$ are the total PBE-D2 energies of the slab and bulk model, respectively, $m$ is the ratio between the slab and bulk stoichiometries, and $A$ is the slab unit area. Factor 2 accounts for both the top and bottom ends of the slab models.

### 3.2 Dynamic simulations

We performed ab initio molecular dynamics (AIMD) simulations, using the CP2K program that employs Gaussian-type and plane-wave basis sets for the wavefunction and the electron density, respectively (Kohlmeier et al. 2013). We run AIMD simulations in order to sample the complex conformational space with the aim to amorphize CO/H$_2$O ice mixtures. All AIMD simulations are performed at the PBE level, enriched by the D3 a posteriori Grimme correction for dispersive interactions (Grimme 2010). A molecular optimized double-$\zeta$ basis set was applied to all atoms together.

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**References**

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- Noël et al. (2001)
- Al-Halabi, van Dishoeck & Kroeis (2004a, 2004b)
- Manca, Roubin & Martin (2000)
- Martin, Manca & Roubin (2000)
- Zicovich-Wilson et al. (2004)
with the corresponding Goedecker–Teter–Hutter (pseudopotential (Vandevondele et al. 2005). The parameters used for the AImD simulations are 0.5 fs for the time-step, 15 ps for the production time, and Nose–Hoover thermostat at 300 K and 1 bar for the pressure.

4 RESULTS AND DISCUSSION

In this section, we will discuss the structural, energetic, and spectroscopic features of the different CO/H\(_2\)O ice systems we simulated in this work.

4.1 Ice models

We are well aware that interstellar ices are likely to be highly amorphous and porous. However, simulating amorphous systems is a non-trivial and highly subjective operation from a computational point of view. Keeping this in mind, we adopted two different 3D periodic models of ice to simulate crystalline ice phases, stable at low temperature and pressure (i.e. in ISM-like conditions): (i) the ordinary hexagonal ice Ih and (ii) the orthorhombic proton-ordered polar ice XI (space group \(\text{Cmc}_2\_1\)) or C-\(\text{ICE}\). In turn, we modelled Ih ice with periodic (long-range ordered) structures belonging to the \(\text{Pna}_2\_1\) space group, usually referred as P-\(\text{ICE}\), with a unit cell providing a variety of configurations to describe the local proton disorder (Pisani, Casassa & Ugliengo 1996; Erba et al. 2009).

Specifically, we selected P-\(\text{ICE}\) to model ice surfaces, by means of 2D periodic slabs, because it has local disorder, yet forms a simple (001) and (010) surfaces with no dipole component perpendicular to them, and essentially no dipole component parallel to them either (Pisani et al. 1996; Casassa & Pisani 2002).

Bulk P-\(\text{ICE}\) was also used to simulate cages (as a single and \(2\times2\times2\) supercell) as well as to simulate a proper mixed ice system (\(1\times1\times2\) supercell). Cages were also simulated by means of clathrate structures (namely, sI and sII) and of a \(2\times2\times2\) C-\(\text{ICE}\) supercell.

4.2 Summary of CO/H\(_2\)O ice models

The CO/H\(_2\)O ice interaction depends on the electronic properties of the CO molecule, such as its electrostatic potential (ESP) and its dipole and quadrupole moment. In Section 4.3, we provide a detailed analysis of these properties, together with those for crystalline solid CO systems. The latter can be representative of situations where ‘solid CO islands’ grow, under interstellar conditions, upon the dust grain mantles.

In Section 4.4, we describe the single adsorption of CO on exposed \(\text{OH}\) sites of two selected P-\(\text{ICE}\) surfaces (\(\text{OH}-\text{CO/OC}\) models). These models can be seen as references for the pure CO/H\(_2\)O ice interaction, since no other factors such as CO/CO repulsion occur. However, they are probably not proper representative of dirty ices grown under interstellar conditions because of the very low CO/H\(_2\)O molecular ratio.

To improve our description, in Section 4.5 we deal with multiple CO adsorption on the (001) P-\(\text{ICE}\) surface. These models can be representative of the formation of a solid-H\(_2\)O/solid-CO interface.

Since interstellar ices are expected to be highly porous (Al-Halabi et al. 2004b), in Section 4.6 we explore the encagagement of one or more CO molecules entrapped inside 3D periodic models of water cages of different size and shape, from the quite large cavities in clathrate named as sI and sII to the very small ones in P- and C-\(\text{ICE}\) bulks. For a complete description, we developed two 2D slab models where one CO is entrapped within ‘surface water cages’, i.e. cages formed by removing some water molecules at the surface of the water ice model (‘expulsion’ models). We also studied CO substituting a H\(_2\)O molecule in the ice bulk (‘Sub\(_\text{CO}\)’ model); calculations revealed, however, a restructuring of the CO surrounding water molecule in a clathrate-like fashion.

Finally, in Section 4.7 we simulate the spectroscopic and energetic properties of a mixed CO/H\(_2\)O ice system with a CO/H\(_2\)O molecular ratio close to the interstellar relative abundances in the solid phase.

4.3 Pure CO

Gas phase carbon monoxide possesses a small dipole moment \(\mu_E\) of 0.122 D oriented as \(\text{C(}\delta^-\text{)}\rightarrow\text{O(}\delta^+\text{)}\) (Muenter 1975). It is well known (Scuseria 1991) that Hartee-Fock and several density functional theory (DFT) methods fail in the prediction of both the absolute value and the orientation of the CO dipole. However, within our adopted computational scheme, CO dipole moment has the correct orientation, even if the module is slightly overestimated (0.196 D). From an electrostatic point of view, CO must be considered as a quadrupolar molecule due to its large quadrupole moment (Graham, Imrie & Raab 1998). In a recent work (Zamirri et al. 2017), we suggested that the quadrupole moment plays a fundamental role in determining the energetic and spectroscopic features of CO adsorbed on selected surfaces of crystalline forsterite as a model of the core dust grain. This quantity, not directly computed by the CRYSTAL code, can be eventually inferred by mapping the molecular ESP superimposed to the electron density. As it is shown in Fig. 1, the quadrupolar distribution of electric charges for gas-phase CO is clearly visible from its ESP map. The value of the \(zz\) component of the quadrupole, computed at the same level of theory of CRYSTAL17 calculations by means of the Gaussian09 code (Frisch et al. 2013), is \(-1.4942\) D Å to be compared with the experimental value of \(-2.63\) D Å (Chetty & Couling 2011).

In order to underline the importance of quadrupolar effects, we run very simple molecular calculations on different CO dimers. In Fig. 1, we report the structures corresponding to the only three energy minima we found. Regardless the initial geometry, in the final optimized ones the two CO molecules assumed an antiparallel configuration, where the electric negative poles of one molecule (C or O atoms) are interacting with the ‘positive belt region’ of the other molecule, resembling the slipped pair in the benzene dimer, entirely dominated by quadrupole–quadrupole interactions. Initial geometries like \(T-\) or \(L\)-shaped ones result in saddle points of first and/or second order. The structures for all tested cases, as well as their spectroscopic features, are reported in the ‘Dipolar and quadrupolar interactions’ section of Appendix A (Figs A13–A19 and Table A3). The C–\(\pi\) interaction results to be almost 50 per cent more stable than the other two (O–\(\pi\) and \(C–\pi\)O–\(\pi\)-\(\text{\_I}\)); see \(\Delta H\) column of Table 1. Our results, although in contrast with those obtained in a previous work (Collings, Dever & McCoustra 2014) at a lower level of theory, are in complete agreement with the fact that the C atoms represent the negative pole of the CO-dipole.

The other extreme is represented by solid \(\alpha\)-CO cubic crystal structure (Vegard 1930) – space group \(P2_13\) – with four molecules per unit cell located on equivalent sites of symmetry \(C_3\). In the equilibrium configuration, the centres of the molecules are at the corners and face centres of a cube and the internuclear axes are oriented along the body diagonals, as in Fig. 1.

Our optimized structure has a sublimation enthalpy at 0 K per CO \(\Delta H\) of 8.0 kJ mol\(^{-1}\) (computed according to Equation(2), including all ZPE contributions in the solid phase), This value is in
Figure 1. Top: molecular electrostatic potential maps superimposed to the electron densities for the single CO molecule and the three energy minima of the CO dimer. The iso-density value for the electron density is set equal to $10^{-5}$ au. $\mu_E$ is the CO electric dipole moment vector. Bottom: optimized structures for solid $\alpha$-CO with $P2_13$ (left) and $P1$ (right) symmetry. Sublimation enthalpies $\Delta H$ in kJ mol$^{-1}$, vibrational shifts $\Delta \bar{\nu}$ in cm$^{-1}$.

Table 1. Structural, energetic, and vibrational features for the three CO dimers of Fig. 1. Interaction energies $\Delta E^{\text{int}}$ and enthalpies $\Delta H$ per CO molecule in kJ mol$^{-1}$, wavenumbers $\nu$ and shifts $\Delta \bar{\nu}$ with respect to the gas-phase CO stretching in cm$^{-1}$, distances in Å.

|        | $\Delta E^{\text{int}}$ (kJ mol$^{-1}$) | $\Delta H$ (kJ mol$^{-1}$) | $d_{C-O}$ (Å) | $d_{X-X}$ (Å) | $\nu$ (cm$^{-1}$) |
|--------|----------------------------------------|---------------------------|---------------|---------------|-----------------|
| C–π    | -1.4 (-0.7)                            | -0.7                      | 1.1367        | 3.249         | 2143 (0)        |
| O–π    | -0.9 (-0.3)                            | -0.5                      | 1.1368        | 3.177         | 2142 (-1)       |
| C–π/O–π| -1.0 (-0.3)                            | -0.5                      | 1.1368        | 3.311         | 2142 (-1)       |

Note: $^a$Without the dispersion contributions in parenthesis

Table 2. Comparison between experimental and our computed data for crystalline solid $\alpha$-CO ($P2_13$ symmetry). $a_0$ lattice parameter in Å, $\Delta H$ sublimation enthalpy in kJ mol$^{-1}$, MIS ‘mean internuclear separation’ (average of the C–O bond lengths) in Å and $\bar{\nu}$ CO stretching wavenumber in cm$^{-1}$.

| Quantity | Experimental | Our calculations | % Error |
|----------|--------------|------------------|---------|
| $a_0$    | $^b$5.64     | 5.54             | 1.8     |
| $\Delta H$ | $^b$8.3     | 8.0              | 3.8     |
| MIS      | $^c$1.128    | 1.136            | 0.7     |
| $\bar{\nu}$ | $^d$2138.1  | 2142.0           | 0.2     |

Notes: $^a$Hall & James (1976).

$^b$Kohin (1960). Measured at 10 K.

$^c$Hall & James (1976).

$^d$Ewing & Pimentel (1961).

These small differences can be explained by different factors. First, experimental IR spectra for solid CO show a broadening in the band, centred at 2138 cm$^{-1}$, as a consequence of the several different contributions in a poly-crystalline-like phase, whereas our computed values refer to a perfect single crystal. Secondly, when scaling all computed wavenumbers for the same factor (computed from gas-phase data) we are assuming that all the perturbing factors in gas-phase (anharmonicity, method-related errors) are implicitly preserved in the solid phase. However, the absolute differences are almost negligible and, within all these assumptions, we are able to reproduce the most important experimental aspect,
Table 3. IR signals, in cm\(^{-1}\), for pure CO systems. Experimental data were recorded at \(T \approx 10\,\text{K}\).

| Phase                   | Experimental | Our calculations |
|-------------------------|--------------|------------------|
| gas                     | \(^{a}\)2143 | 2143             |
| liquid                  | \(^{b}\)2138 | –                |
| bulk (\(\alpha\) phase) | \(^{c}\)2138 | 2142             |
| bulk \(\alpha\)\(^{13}\)CO | \(^{d}\)2092 | 2094             |
| bulk LO/TO              | \([2143.7, 2138.5]\) | [2147, 2142]     |

Notes: \(^{a}\)Mina-Camilde et al. (1996).  
\(^{b}\)Ewing (1962).  
\(^{c}\)Ewing & Pimentel (1961).  
\(^{d}\)Pontoppidan et al. (2003).

i.e. the significant redshift of solid CO with respect to the gas phase.

Finally, to enlighten on the LO/TO splitting origin of the peak at 2143.7 cm\(^{-1}\), as suggested by Pontoppidan et al. (Pontoppidan et al. 2003; Collings et al. 2003b), we have evaluated the CO dielectric constant \(\varepsilon = 1.69\) and calculated the LO wavenumber. Again, the computed value is overestimated with respect to the experimental signal (see Table 3), but the interval shift of \(\approx 5\,\text{cm}^{-1}\) between the LO and TO modes is perfectly reproduced. On this basis, the LO/TO hypothesis appears consistent although it cannot be observed in astronomical spectra (Collings et al. 2003a) and Boogert and coworkers (Boogert et al. 2015) have recently found this assignment less likely.

4.4 Single CO adsorption on water ice surfaces

Single CO adsorption has been studied on two different ice surfaces cut out from P-ICE bulk, corresponding to the (001) and (010) faces. The 2D slabs, consisting of 12 water molecules per unit cell, as shown in Figs 2 and 3, were fully optimized and results show that the (010) is slightly more stable than the (001) one (the former having a \(E_S\) value 3.6 mJ m\(^{-2}\) lower than that for the latter). Both surfaces expose in equal proportion dangling \(\text{D}_\text{H}\) and \(\text{D}_\text{O}\) atoms.

One CO molecule per unit cell was approached to the (010) and (001) \(\text{D}_\text{H}\) surface sites through its C and O atoms. These configurations, reported in Figs 2 and 3 for the (010) and (001) cases and described in Table 4, will be referred in the following as \(\text{D}_\text{H} \cdot \cdot \cdot \text{CO}\) and \(\text{D}_\text{H} \cdot \cdot \cdot \text{OC}\), respectively. On both surfaces, the initial \(\text{D}_\text{H} \cdot \cdot \cdot \text{CO}\) interaction is preserved while the \(\text{D}_\text{H} \cdot \cdot \cdot \text{OC}\) is lost during the optimization as dispersive and quadrupolar interactions between CO and the surface dominate over H-bonds. In all cases, CO quadrupolar moment exerts a crucial effect on both energetic and spectroscopic features. Irrespective of the surface, there are no evidences of a direct interaction involving \(\text{D}_\text{O}\) sites, neither via the O atom (positive pole of the dipole) nor via the \(\pi\) system (positive region of the quadrupole). This trend is also confirmed by calculations performed on a single CO interacting with the water-trimer, which represents the simplest system exposing both \(\text{D}_\text{H}\) and \(\text{D}_\text{O}\). Structural, energetic, and vibrational data for these systems are reported in Appendix A (Figs A1–A12 and Table A3).

For \(\text{DH} \cdot \cdot \cdot \text{CO}\) cases, a definitive H-bond is formed between the C atom and the \(\text{D}_\text{H}\) on both surfaces, resulting in a blueshifting of \(\approx 15\,\text{cm}^{-1}\) of the CO stretching wavenumber with respect to the gas phase and in a shortening of the C–O bond. These results are consistent with those obtained by Allouche and coworkers (Allouche...
Figure 3. Final optimized geometries for CO initially interacting with nH sites of the (001) P-ICE surface via either C (nH·CO) or O (nH·OC) atom. Top (left) and side views along a (centre) and b (right) periodic vectors. z is the direction perpendicular to the cut. In top views, C and O atoms of CO molecules as van der Waals spheres, nHs in dark grey.

Table 4. Different configurations of CO interacting with perfect models of P-ICE surfaces. Interaction energies $\Delta E^\sigma$ and enthalpies $\Delta H$ per CO molecule in kJ mol$^{-1}$ ($\Delta E^\sigma$ without the dispersion contribution in parenthesis), wavenumbers $\bar{\nu}$ and shifts $\Delta \bar{\nu}$ with respect to the gas-phase CO stretching in cm$^{-1}$, C–O distances $d_{C-O}$ and H-bonds in Å.

| Configuration          | $\Delta E^\sigma$ (no disp.) | $\Delta H$ | $\bar{\nu}$ ($\Delta \bar{\nu}$) | $d_{C-O}$ | H-bond |
|------------------------|-------------------------------|------------|-----------------------------------|-----------|--------|
| CO adsorbed on P-ICE   |                               |            |                                   |           |        |
| (010) surface          |                               |            |                                   |           |        |
| nH–CO                  | $-21.0$ ($-12.9$)             | $-18.4$    | 2158 (+15)                        | 1.1351    | 2.215  |
| nH–OC                  | $-14.7$ ($-3.3$)              | $-12.9$    | 2114 ($-29$)                      | 1.1402    |        |
| (001) surface          |                               |            |                                   |           |        |
| nH–CO                  | $-19.1$ ($-9.9$)              | $-16.8$    | 2163 (+20)                        | 1.1345    | 2.375  |
| nH–OC                  | $-15.3$ ($-5.4$)              | $-13.3$    | 2122 ($-21$)                      | 1.1393    |        |
| CO/P-ICE interfaces    |                               |            |                                   |           |        |
| Interface 1            | $-5.1$ ($-2.9$)               | $-4.9$     | 2133 ($-10$)                      | 1.1377    |        |
|                        |                               |            | 2140 ($-3$)                       | 1.1369    |        |
|                        |                               |            | 2147 (+4)                         | 1.1360    |        |
|                        |                               |            | 2164 (+21)                        | 1.1344    | 2.315  |
| Interface 2            | $-5.1$ ($-3.0$)               | $-5.0$     | 2137 ($-6$)                       | 1.1372    |        |
|                        |                               |            | 2145 (+2)                         | 1.1362    |        |
|                        |                               |            | 2150 (+7)                         | 1.1361    |        |
|                        |                               |            | 2162 (+19)                        | 1.1347    | 2.290  |

et al. 1998; Manca et al. 2001) and further confirmed by Al-Halabi et al. (2004a) via molecular dynamics simulations. For nH·OC cases, specific dispersive interactions cause an increasing of the C–O bond and a consequent redshift of $\approx -20$ cm$^{-1}$. Summarizing the results of this subsection, we can state that (i) the nH·CO interactions may explain the experimentally observed blueshifted peak at around 2152 cm$^{-1}$; (ii) there is no interaction between CO molecules and nO sites; (iii) quadrupolar and dispersive interactions play a fundamental role in determining the energetic and structural features of CO adsorption, particularly when CO interacts via its O atom with nH sites; and (iv) the resulting CO redshifted band does not provide a convincing explanation for the experimental signal at 2138 cm$^{-1}$.

4.5 CO/H$_2$O ice interface: multiple CO adsorption

As the concentration of CO increases up to the monolayer limit, a different kind of model is needed to reproduce the CO/H$_2$O ice interaction. Then, we have designed two different structures by
placing four CO molecules per unit cell upon the (001) surface of P-ice, as in the initial panels of Fig. 4. The optimization procedure yields an interesting rearrangement in both CO/H$_2$O ice interfaces. As reported in the final panels of Fig. 4, CO molecules reorganize to maximize the reciprocal quadrupolar interactions and the H-bonds with the surface. The resulting structures are two CO-layered slabs, resembling the arrangement of CO molecules in the bulk structure. Energetics, computed according to Equations (1) and (2), and spectroscopic features are reported in Table 4.

As a consequence of the different surrounding of each CO molecule, the CO stretching band splits in few components and the overall features of the computed spectra are the same for the two interfaces. In particular, adopting the CO labels of Fig. 4, it can be noted that the high wavenumber peaks, around 2160 cm$^{-1}$, correspond to the interaction of the carbon atom of the CO$_A$ molecule with the OH of the water slabs, as in the case of the OH-CO interaction discussed in Section 4.4. Similarly to OH-OC, the most redshifted feature is due to the stretching of the CO$_D$ molecules, which are those exposing their O atoms to the H$_2$O ice surface. The two remaining vibrational modes, which involve the CO molecules of the upper layer, are evidently less affected by the icy water surface and reflect the CO-CO interactions.

Therefore, it can be stated that, as the availability of OH sites is over, the exceeding CO molecules reorganize and grow up as homogeneous layers on top of the ice surface. As a consequence, a broadening around the 2138 cm$^{-1}$ IR signal takes place due to the different interactions of the adsorbed CO molecules, which leave red- and blueshifts.

In very recent works, the ‘spontlectric’ nature of CO layers grown on water ices has been revealed. It has been suggested that this peculiar behaviour could strongly affect the physico-chemical evolution of clouds since it could influence the abundances of charged chemical species in the gas-phase ISM (Lasne et al. 2015; Rosu-Finsen et al. 2016). We hope that our interface models could be useful to computationally highlight this important aspect in further studies.

4.6 CO in H$_2$O ice cages

Sandford et al. (1988) observed that as the amorphous H$_2$O ice is warmed up, from 10 to 150 K, the H$_2$O molecules rearrange into a more ordered structure. Moreover, the annealing process seems, on the one hand, to facilitate CO diffusion into ice pores (Palumbo 1997), as attested by the narrowing of the signal at 2136 cm$^{-1}$ and, on the other hand, it is claimed to be responsible for the appearance of new adsorption features around 2144 cm$^{-1}$ (4.664 µm) that could be possibly related to the formation of clathrate-like structures (Sandford et al. 1988).

In order to investigate on these aspects, we simulated the CO encagement into different H$_2$O-crystalline lattices ranging from ordinary ice to some clathrates of increasing cavity volume.

In particular, with reference to Fig. 5, we addressed the natural occurring type I CO-clathrate (Davidson et al. 1987; Sloan & Koh 2007), sI, whose two cages were filled with one CO molecules per cage, and the recently discovered structure III methane–clathrate (Loveday et al. 2001), sIII which, despite its similarity with hexagonal ice, shows larger cavities. Here we inserted two CO molecules.
Figure 5. CO entrapped in different cage models. Enthalpy variations per CO molecule $\Delta H$ in kJ mol$^{-1}$, vibrational shifts $\Delta \bar{\nu}$ in cm$^{-1}$, distances in Å.
Table 5. Energetic, spectroscopic, and structural data for CO entrapped in different water–ice cages. Interaction energies $\Delta E^\circ$ and enthalpies $\Delta H$ per CO molecule in kJ mol$^{-1}$ ($\Delta E^\circ$ without the dispersion contribution in parenthesis), wavenumbers $\bar{\nu}$ and shifts $\Delta \bar{\nu}$ with respect to the gas-phase CO stretching in cm$^{-1}$, C–O bond lengths $d_{C-O}$ in Å. $\Delta V$ is the percentage variation in cell volume with respect to the initial structure. smI cases can be divided into two different groups according to their spectroscopic features (see the text). ‘Subco’ stands for substitutional CO (see the text and Fig. 5). $\text{smI}_{\text{II}}$ refers to only one CO molecule per cage in sm structure.

| Model       | $\Delta E^\circ$ (no disp.) | $\Delta H$ | $\% \Delta V$ | $\bar{\nu}$ ($\Delta \bar{\nu}$) | $d_{C-O}$ |
|-------------|-----------------------------|------------|----------------|-----------------------------------|-----------|
| sl          | $-16.1 (+4.8)$              | $-14.5$    | $-2139 (-4)$   | 1.1369                            |           |
| $\text{slII}_{\text{I}}$ | $-13.6 (+9.0)$             | $-11.3$    | $+12133 (-10)$ | 1.1372                            |           |
| smI_{0}     | $-7.7 (+18.2)$              | $-4.6$     | $+12137 (-6)$  | $+11362$                          |           |
|             |                             |            |                | 2142 (-1)                         | 11372     |
|             |                             |            |                | 2142 (-1)                         | 11372     |
|             |                             |            |                | 2145 (+2)                         | 11372     |
| $\text{smI}_{\text{II}}$ | $-7.6 (+17.8)$             | $-4.5$     | $+12144 (+1)$  | $+11357$                          |           |
|             |                             |            |                | 2145 (+2)                         | 11372     |
|             |                             |            |                | 2147 (+4)                         | 11372     |
|             |                             |            |                | 2148 (+5)                         | 11372     |
| P-ICE SC    | $+56.0 (+83.1)$             | $+65.6$    | $+21313 (-12)$ | 1.1369                            |           |
| Expulsion 1 | $-10.9 (+31.1)$             | $-9.7$     | $-2140 (-3)$   | 1.1368                            |           |
| Expulsion 2 | $-7.7 (+13.5)$              | $-5.4$     | $-2139 (-4)$   | 1.1367                            |           |
| Subco       | $-3.8 (+18.8)$              | $-0.4$     | $-2134 (-9)$   | 1.1373                            |           |

Note: *All CO molecules have the same $d_{C-O}$ distance.

per cage defining eight different initial configurations and exploiting symmetry. According to their spectroscopic features, these eight initial configurations can be grouped into two groups, labelled as $\text{smI}_{\text{II}1}$ and $\text{smI}_{\text{II}2}$. We also considered a single CO occupancy, $\text{slII}_{\text{I}}$. Finally, the eventuality of CO encagement into crystalline ice was simulated by a $2 \times 2 \times 2$ supercell models (SC) of both C- and P-ICE. All the systems were fully optimized and the vibrational properties for the CO molecules are summarized in Table 5.

From our results, some general considerations arise. Data reported in Table 5 underline the fundamental role of dispersive interactions in determining the energetics of the encagement process. Indeed, there is no evidence of H-bonds involving the C atom of CO molecules with any water molecule of the surrounding so that without the dispersion contribution CO will be unbound. Nevertheless, in the cases of C- and P-ICE SC steric hindrance predominates also when accounting for dispersive forces, i.e. the 12-H$_2$O molecules cavity is too small to host carbon monoxide, as indicated by the positive interaction energies. Secondly, COs induce a small but sensitive relaxation of the water lattice, resulting in an increase of the cell volume between 1 per cent to 10 per cent with respect to the empty initial structures in all cases but sl.

The analysis of spectroscopic data shows that the formation of different clathrate-like cages around CO molecules can partially explain the observed IR signals for CO in water-dominated environments. The redshift of the CO stretching wavenumber with respect to the gas phase is significantly less than that in the case of DH–OC. This behaviour can again be explained by invoking quadrupolar interactions between the ice lattice and the CO. The extent of the redshifting effect for those cases where a single cage hosts a single CO molecule (i.e. sl, $\text{slII}_{\text{I}}$, Subco, C- and P-ICE SC) seems to suggest a general trend between volume of the cage and the redshift itself. However, the volume of the cage cannot be properly defined and thus we performed a detailed neighbouring analysis for the C and O atoms of COs. All data are reported in the ‘Neighbourhood analysis for CO in water cages’ section of Appendix A (Table A6). Briefly, results indicate that those cases characterized by very small cages (as P- and C-ICE SCs) have very positive interaction energies and present the most redshifted wavenumbers, while when the dimensions of the cages are larger (as in the sm case), the entity of the redshift is small.

A noticeable difference appears in the eight sm configurations explored. Despite their similarity in terms of CO–CO and CO–lattice distances (see Fig. 5), in the $\text{slII}_{\text{II}}$ all the four CO stretching modes are blueshifted, while in the $\text{smI}_{\text{II}}$, only one mode is blueshifted.

To enlighten on this aspect, we removed the symmetry and reoptimized the $\text{smI}_{\text{II}}$ and $\text{smII}_{\text{II}}$ structures in P1 space group, ending up with the same frequency pattern. For a more detailed analysis of the IR signals with and without symmetry, see Table A5 of Appendix A. We conclude that these blueshifted modes are not a symmetry artifact but the result of the peculiar geometry assumed by the CO molecules inside the cages giving rise to specific intermolecular interactions. As a matter of fact, these signals are completely absent in the single occupied $\text{slII}_{\text{I}}$ system.

Our data suggest that both the astronomical features of CO/H$_2$O mixtures and the controversial redshifted peaks in the experimental spectra could be due to trapped CO experiencing the potential field inside a water cage and/or via quadrupolar interactions of neighbouring CO molecules.

On the other hand, the very high energetic data for CO entrapped in crystalline C- and P-ICE suggest that carbon monoxide could not be able to diffuse into such small pores but does prefer to remain adsorbed onto proper surface moieties or eventually to rearrange the network of water molecules around it.

In order to test this possibility, from the 3D periodic structure of the P-ICE SC of Fig. 5, we cut two different slab models along the [010] direction so that the water cages hosting COs result to be exposed to vacuum. Surfaces are more flexible structures than bulk so they represent natural escaping routes. The initial and final geometries of these two models are shown in Fig. 6, while the computed properties are reported in Table 5, ‘Expulsion’ rows. Once again, the role of dispersive interactions is fundamental to determine the final energetics and structural features. CO molecules tend to be expelled from the surface pores and only weak quadrupolar interactions keep them inside the structures. However, these ‘expulsions’ cause a great deformation of the surrounding lattice and neighbour water molecules rearrange into ‘semicages’ partially restoring clathrate-like cavities.

These simulations suggest a ‘hydrophobic’ behaviour of CO and clearly show that CO would hardly diffuse into water–ice pores unless they are at least as large as in clathrates. This behaviour is confirmed by the ‘substitutional CO’ case (Subco in Fig. 5 and Table 5), where one water molecule of the P-ICE bulk structures is replaced with one CO interacting, via H-bonds, with two H$_2$O molecules. Once the structure is relaxed, water lattice re-arrange significantly, CO–H$_2$O directional bonds are lost and a clathrate-like cage forms around the carbon monoxide (compare initial and final Subco panels of Fig. 5). The energy balance, reported in the last row of Table 5, takes into account both the deformation energy and the formation of new H-bonds between H$_2$O molecules and results in an almost zero value. This result suggests that the formation of a CO/H$_2$O mixed lattice is rather unlikely, at least at low concentration of carbon monoxide and at the very low temperature characterizing the colder regions of the ISM.
Figure 6. Side views for the 'expulsion' models, i.e. CO expelled from superficial P-ICE pores. Initial geometries on the left, final ones on the right. Enthalpy variations per CO molecule $\Delta H$ in kJ mol$^{-1}$, vibrational shifts $\Delta \nu$ in cm$^{-1}$, C–O distances in Å.
Table 6. Structural, energetic, and spectroscopic features for the five snapshots selected from the AIMD run for the mixed CO/H$_2$O ice system. AIMD times in ps, energetic quantities per CO in kJ mol$^{-1}$, wavenumbers and shifts in cm$^{-1}$, distances in Å. For the 15.0 ps, we computed the energetic features for all the six CO molecules, individually (see the text).

| MD time | $\Delta E^{\text{CP}}$ (no disp.) | $\Delta H$ | $\bar{\nu}$ ($\Delta \bar{\nu}$) | $d_{\text{CO}}$ | H-bond |
|---------|----------------------------------|-----------|---------------------------------|---------------|--------|
| 5.0     | $-2.2$ ($+13.7$)                 | $+0.5$    | $2217 (-26)$; $2148 (+5)$       | $1.1392; 1.1359$ | 2.021, 2.394 |
| 7.5     | $-10.6$ ($+3.4$)                 | $-7.7$    | $2120 (-23)$; $2174 (+31)$      | $1.1390; 1.1330$ | 2.006, 2.207, 2.460 |
| 10.0    | $-6.5$ ($+9.9$)                  | $-3.4$    | $2123 (-20)$; $2177 (+34)$      | $1.1386; 1.1325$ | 1.937, 2.150 |
| 12.5    | $-7.3$ ($+7.2$)                  | $-7.1$    | $2135 (-8)$; $2175 (+32)$       | $1.1371; 1.1329$ | 2.013, 2.078 |
| 15.0    | $-9.5$ ($+6.6$)                  | $-6.3$    |                                 |               |        |
| A       | $-25.9$ (0.0)                    | $-21.7$   | $2172 (+29)$                     | $1.1333$       | 2.036  |
| B       | $-23.9$ ($-2.8$)                 | $-20.3$   | $2164 (+21)$                     | $1.1341$       | 2.055  |
| C       | $-12.8$ ($+11.1$)                | $-10.1$   | $2138 (-5)$                      | $1.1369$       | –      |
| D       | $-15.5$ ($+6.8$)                 | $-12.9$   | $2135 (-8)$                      | $1.1374$       | –      |
| E       | $-18.8$ ($+6.9$)                 | $-15.8$   | $2131 (-12)$                     | $1.1376$       | –      |
| F       | $-5.6$ ($23.8$)                  | $-1.8$    | $2130 (-13)$                     | $1.1374$       | –      |

Figure 7. Left: view along the $b$ (top) and $a$ (bottom) lattice vector for the mixed CO/H$_2$O system. C and O atoms of CO molecules as van der Waals spheres. Right: detailed views of the local water environments surrounding CO molecules. CO molecules labelled from A to F according to their stretching wavenumbers. Enthalpy variation per CO molecule $\Delta H$ in kJ mol$^{-1}$, vibrational shifts $\Delta \bar{\nu}$ in cm$^{-1}$.

4.7 CO entrapped in water-rich mixture

Sub$_{\text{CO}}$ can be considered as the simplest model of a CO/H$_2$O mixed ice. In order to improve the description of these binary systems, we defined a more realistic model by replacing 6 out of 24 water molecules by 6 CO molecules in a $1 \times 1 \times 2$ supercell of P-ICE to reproduce the CO:H$_2$O ratio usually reported for the ISM (Boogert et al. 2015). The initial geometry was amorphized by running an AIMD simulation with the CP2K code at 300 K for 15.0 ps. After that, five snapshots were selected, corresponding to 5.0, 7.5, 10.0, 12.5, and 15.0 ps of production and each structure was optimized at PBE-D2 level with the CRYSTAL17 code. On the final optimized geometries, a complete energy and frequency analysis was performed. Data reported in Table 6 confirm once again the fundamental role of dispersive interactions to define the energetic configuration of these systems where, as in the case of Sub$_{\text{CO}}$, clathrate-like structures build up around CO molecules.

The 15.0 ps optimized structure (see Fig. 7) turns out to be a reliable model of a possible structure since its spectrum closely
Figure 8. Simulated IR spectra for the different AIMD snapshots as a function of the AIMD production time. The values of the vibrational shifts (in cm$^{-1}$) are reported close to all distinguishable peaks. The FWHM for each component of the different spectra is set equal to 15 cm$^{-1}$.

resembles the experimental ones; therefore, we performed a detailed energetic analysis to determine the interaction energies and enthalpies of each individual CO molecule within the mixed CO-ice lattice (Table 6). These quantities were estimated according to Equations (1) and (2) by setting to zero the deformation energy of the mixed [18 H$_2$O + 5 CO] system.

As in the SubCO case, the water lattice tends to form cage structures closely resembling the natural cavities in clathrates. Thus, CO molecules entrapped within these cages, and not involved in H-bonds, are characterized by redshifts. However, during the evolution of the simulation, H-bonds among CO and water can actually form in our model. The CO molecules involved in these H-bonds have, in accordance with our other results on P-ICE slabs, quite large blueshifts. The absolute computed IR intensities for this blueshifted signals are smaller than those for redshifted ones. The lack of the 2152 cm$^{-1}$ peak, which, together with our other previous results, was assigned to the DH··CO interaction, might be explained by a low concentration of DH sites in water-dominated ices grown under interstellar conditions.

Referring to Fig. 7 and using the adopted notation, it can be seen that the A and B CO molecules, involved in H-bonds, have the highest strength of interaction whereas for all the other COs the energetic is consistent with dispersion-driven interactions (Table 6). It is worth noting that the surroundings of A and B CO molecules are very different from the DH··CO ones as occurring on P-ICE surfaces since these molecules are entrapped in water cages.

To appreciate the evolution of the spectroscopic signals, the five IR spectra of the corresponding AIMD snapshots are superimposed in Fig. 8. All spectra except the 7.5 ps one present a main persistent component at $\approx -7$ cm$^{-1}$. A more redshifted feature at $\approx -20$ cm$^{-1}$ rapidly disappears, while after 7.5 ps, two blueshifted signals appear at around $+20$ and $+30$ cm$^{-1}$, but while the former is preserved in the final spectrum, the latter also disappears, resulting in a shoulder of the main blueshifted peak.

4.8 Merging the IR signals: global simulated spectrum

As a general observation, the computed absolute IR intensities for the CO blueshifted frequencies are lower than those for redshifted ones (see Tables 7 and A4). This is the result of an 'enhancement process' that can be underlined by comparing the absolute intensities of a single CO interacting via its C and O atoms with the H atom of a water molecule (Fig. A20 of Appendix A). In the first case (H-bond with the C atom), the CO stretching wavenumber is blueshifted by 16 cm$^{-1}$ and has an absolute intensity of 60.86 km mol$^{-1}$ (see HOH··CO line in Table 7), slightly lower than that for gas-phase CO (62.04 km mol$^{-1}$). On the other hand, when an H-bond forms between water and the O atom of the CO molecule, the CO stretching wavenumber is redshifted by 7 cm$^{-1}$ and the absolute intensity is enhanced by more than 25 per cent (78.28 km mol$^{-1}$) with respect to the gas-phase (HOH··OC line in Table 7). This is in accordance with previous results where a similar enhancement has been observed for CO interacting with selected metal ions (Ferrari, Ugliengo & Garrone 1996). This enhancement of redshifted wavenumbers could partially explain the lack of the blueshifted component in astronomical spectra that is very faint and partially hidden.
Table 7. Resume of our energetic and spectroscopic computed data. Enthalpy variations per CO $\Delta H$ in kJ mol$^{-1}$, IR stretching wavenumbers $\nu$ and shifts $\Delta \nu$ in cm$^{-1}$. Absolute IR intensities (in km mol$^{-1}$) per CO are also reported.

| Model                  | Figure | $\Delta H$ | $\nu (\Delta \nu)$ | Intensity | Model | Figure | $\Delta H$ | $\nu (\Delta \nu)$ | Intensity |
|------------------------|--------|------------|---------------------|-----------|--------|--------|------------|---------------------|-----------|
| Gas phase CO           | 1      | 2143 (0)   | 62.04               |           | 5      | 2139 (−4) | 23.18       |         |           |
| HOH-CO                 | A20    | 2159 (+16) | 60.86               |           |        | 2140 (−3) | 181.04      |         |           |
| HOH-OC                 | A20    | 2136 (−7)  | 78.28               |           | 4      | 2137 (−6) | 78.36       |         |           |
| Solid $\alpha$-CO symm.| 1      | 2139 (−4)  | 0.00                |           | 1      | 2142 (−1) | 49.88       |         |           |
| Solid $\alpha$-CO no symm. | 1   | 2135 (−8)  | 110.83              |           | 2      | 2145 (+2)  | 0.00        |         |           |
|                         |        | 2136 (−7)  | 109.40              |           | 3      | 2144 (+1)  | 0.25        |         |           |
|                         |        | 2142 (−1)  | 94.00               |           |        | 2142 (−1) | 0.00        |         |           |
|                         |        | 2145 (+2)  | 0.03                |           |        | 2141 (−4) | 44.62       |         |           |
| P-Icem: (010) nH-CO    | 2      | 2158 (+15) | 52.49               |           | 5      | 2141 (+4)  | 55.94       |         |           |
|                         |        | 2114 (−29) | 92.98               | C-Icem SC | 4      | 2131 (−12) | 135.72      |         |           |
| P-Icem: (001) nH-CO    | 2      | 2163 (+20) | 49.07               |           | 5      | 2131 (−12) | 139.15      |         |           |
|                         | 3      | 2122 (−21) | 90.15               | SubCO     | 6      | 2140 (−3)  | 100.2       |         |           |
| P-Icem: interface 1    | 4      | 2133 (−10) | 106.62              |           | 7      | 2130 (−13) | 179.18      |         |           |
|                         |        | 2140 (−3)  | 70.54               | P-Icem: exp1 | 6     | 2140 (−3)  | 100.2       |         |           |
|                         |        | 2147 (−4)  | 49.41               | P-Icem: exp2 | 6     | 2139 (−4)  | 85.59       |         |           |
|                         |        | 2164 (−21) | 48.88               | Mixed ice: 15.0 ps AIMD | 7      | 2130 (−13) | 179.18      |         |           |
| P-Icem: interface 2    | 4      | 2137 (−6)  | 87.07               |           |        | 2131 (−12) | 66.05       |         |           |
|                         |        | 2145 (+2)  | 81.81               |           |        | 2138 (−5)  | 101.14      |         |           |
|                         |        | 2150 (+7)  | 33.69               |           |        | 1764 (+21) | 97.14       |         |           |
|                         |        | 2162 (+19) | 46.43               |           |        | 2172 (+29) | 67.92       |         |           |
| sIII$_{11}$            | 5      | 2133 (−10) | 95.82               |           |        | 2135 (−8)  | 127.69      |         |           |

Notes: $^a$Sublimation enthalpy. $^b$IR inactive mode.

In Fig. 9 we reported the overall spectra obtained by summing the computed absolute intensities per CO molecule (i.e. the absolute intensity of a specific vibrational mode divided by the number of COs involved in that specific mode) of all tested cases for each category. The FWHMs for each component of the simulated spectra were arbitrarily set equal to 3 cm$^{-1}$; this value is comparable with those reported for the three components of the 4.65 $\mu$m CO band in interstellar ices (Boogert et al. 2015). In the 'cages' panel, we a priori excluded SC models because of their highly unfavourable interaction enthalpies (Tables 5 and 7). For a better comparison, we also highlighted the 2143 cm$^{-1}$ (4.666 $\mu$m) free CO signal (blue continuous line) as well as the 2143.7, 2139.9, and 2136.5 cm$^{-1}$ (4.665, 4.673, and 4.681 $\mu$m) ones corresponding to Ap (CO$_2$/CO > 1 and CO > 90 per cent) and Pr interstellar ices (red dot–dashed, green dashed, and black-dotted lines, respectively; Boogert et al. 2015). However, it is worth underlining the fact that our models cannot be interpreted in terms of Ap and Pr ices, mainly because we did not test proper mixed ices (i.e. also including other species such as CO$_2$, CH$_3$OH, CH$_4$, etc.); indeed, our aim was to elucidate the structural, energetic, and spectroscopic features of CO in different water–ice structural environments. Nevertheless, some conclusions can be drawn as both the APOLAR and POLAR components can result from CO involved in very different physico-chemical situations, from pure CO systems (solid CO panel) to mixed ones (15.0 ps AIMD panel), passing from layered and encaged models. We merged all our results in the single, overall spectrum of Fig. 10, that is, the sum of all panels of Fig. 9 excluding the nHs one. nH cases have been excluded because they envisage slab models that underestimate the typical CO: H$_2$O ratio in the ISM (Boogert et al. 2015). For sake of semiquantitative comparison, the observed spectrum belonging to the LYSO L 1489 IRS object from (Boogert et al. 2015) was superimposed. The prediction of the maximum of the peak is in good agreement with the astronomical and experimental observations, even if the band width is slightly overestimated. Blueshifted peaks are visible in the computed spectrum with intensities definitely lower than the main peak. As we have already addressed, the reason for the missing of these blueshifted signals in astronomical spectra is probably due to the low concentration of OH sites in interstellar ices, together with what previously asserted for the IR intensities of the D+ CO interaction. An interesting and alternative explanation to this problem provides that OH sites could be involved in stronger interactions with other molecular species (Fraser et al. 2004), with particular reference to CO$_2$ (Garrod & Paudy 2011).

From a pure energetic point of view, our results underline how CO is only physisorbed to the water ice matrix. Indeed, the enthalpy values fall in an $\sim$ 5–18 kJ mol$^{-1}$ range (excluding SC models and the Sub$_{CO}$ one where a large part of the interaction energy is consumed in the water lattice rearrangement) and thus weak interactions such as dispersive and electrostatic ones are prominent. Our data can be compared with those usually inferred from classical thermal programmed desorption (TPD) experiments, where different monolayers (MLs) of CO are initially adsorbed onto selected surfaces, such as water–ice ones, and then desorbed. Referring to Smith, May & Kay (2016), the desorption energies for CO adsorbed on amorphous solid water at very low coverages (< 0.05 ML) are slightly greater than 15 kJ mol$^{-1}$, in good with our computed $\Delta H$ values for D+H cases, in particular the D+ CO ones. On the opposite side, the desorption energies for two monolayers can be derived to be approximately equal to 7 kJ mol$^{-1}$, to be compared with our P’ Icem interface’ models ($\sim$ 5 kJ mol$^{-1}$; Tables 4 and 7). However, please note that a proper comparison cannot be performed because of the lack of a generic definition of monolayer, from both an experimental and computational point of view.
IR spectral fingerprint of carbon monoxide in interstellar water–ice models

5 RESUME AND CONCLUSIONS

Adsorption of organic molecules on or inside ice surfaces of the dust mantle of interstellar grains has received significant attention in astrochemical investigations due to their implication in reactions yielding precursor molecules for life. Specifically, from a chemical point of view, the interest around CO is mainly due to the possible formation of derived species such as carbon dioxide, methanol, formaldehyde, and formic acid (Karssemeijer et al. 2014). In the meanwhile, impressive quantity of spectroscopic measurements has become available and several laboratory and theoretical simulations have crossed their intuition to provide reliable molecular configurations to be associated with each spectroscopic peak. As regards CO, the debate is mainly concentrated on the red- and blueshifts of its stretching frequency recorded as ISM signals with respect to its characteristic gas-phase frequency.

Then, in order to enlighten on this very small and informative portion of the spectrum, we simulated a wide range of CO/H$_2$O ice environments adopting different models of water ice (i.e. crystalline and amorphous ice bulk, surfaces, clathrates) by means of DFT simulations (PBE-D2 level of theory) on periodic and molecular models. Our main goal was to furnish solid theoretical justifications on this topic, with particular reference to the interpretation of both the experimental and astronomical spectra. Because the vibrational shifts involved are very small, we ensured that the calculations were run with an accurate level of theory.

The present results, far from being conclusive, provide some general indications on CO/H$_2$O ice binary system, both as regards its properties and spectroscopic fingerprint:

(i) Dispersive and quadrupolar interactions have a prominent role in determining the structural and spectroscopic features of carbon monoxide interacting with other systems, e.g. other CO and/or water molecules. Within our computational methodology, the electrostatic properties of the CO molecule (orientation of the dipole and quadrupole electric moments) are correctly described, even if the value of the CO quadrupole is underestimated with respect to the experimental one.

(ii) The most favourable interaction between carbon monoxide and water involves dangling H atoms (D$_H$) of crystalline H$_2$O ice (P-ICE) surface models and the C atom of CO (negative pole of both
Figure 10. Overall simulated spectrum obtained as a sum of all panels of Fig. 9 except the DHs one. The FWHM of each component is set equal to 3 cm$^{-1}$. In blue, free gas-phase CO stretching reference. The maximum of the signal is at 2136 cm$^{-1}$ (4.682 µm). The solid black line represents the observed spectrum for LYSO L 1489 IRS object. This spectrum is taken and adapted from Boogert et al. (2015).

dipole and quadrupole). It is this specific DH··CO interaction that causes a blueshift of the CO stretching wavenumber with respect to the gas-phase value. The symmetric DH··OC interactions is weaker and is lost during the geometry optimizations – as the coupled action of dispersive and quadrupolar forces overcomes the DH··OC bond, with a corresponding redshift of the C–O stretching frequency. The large shift values (in module) for these reference CO/H$_2$O ice models (in particular the large redshifts), together with the very low CO/H$_2$O ratio, suggest that they are probably not representative of a real astronomical ices, at least considering the need of a crystalline ice.

(iii) The lack of the blueshifted peak in astronomical spectra could be the result of two factors: first, blueshifted signals are intrinsically less intense than redshifted ones and, secondly, DH sites are probably scarcely present (but not completely absent) in the interstellar ices covering dust grain cores, probably due to the amorphous nature of ice in which all available DH are engaged in H-bond to maximize the interaction energy within the ice mantle. Please note that alternative explanations to this problem have also been suggested.

(iv) The redshifted peak around 2138 cm$^{-1}$ present in both astronomical and experimental spectra could be the result of CO involved in different reciprocal configurations with respect to the water–ice matrix such as layered CO/H$_2$O ices, CO entrapped within water cages of different sizes, pure (crystalline-like) solid CO or proper mixed water ices. For all these models, small redshifts appear as a result of the establishment of weak, non-H-bond interactions. However, these interactions could also produce (in all cases but solid CO) small blueshifted signals.

(v) CO shows a 'hydrophobic behaviour', which hinders it to penetrate into small ice pores. As a consequence, the presence of CO can induce large rearrangements of the surrounding lattice water molecules, resulting in the formation of clathrate-like cages where a perfect network of H-bonds among H$_2$O molecules establishes, as happens for the 'expulsion' and 'Sub$_{CO}$' models. As a consequence, CO could not simply substitute water molecules in the ice lattice.

(vi) As the concentration increases, CO molecules tend to clusterize and/or rearrange forming homogeneous structures bonded together by attractive dispersive and CO-quadrupole-dominated interactions, whose morphology is dictated by the ice degree of order. Namely, ordered CO layers form on top of a preformed water–ice surface, whereas CO molecules clusterize inside a proper mixed CO–H$_2$O matrix. Our results indicate that small blue- and redshifted peaks could appear in the IR spectrum as a consequence of this 'clusterization' process.

We hope this work could be useful to enlighten some aspects of this important astronomical topic, and to provide atomistic details of the CO/H$_2$O ice interactions that could be incorporated in further studies more involved in the problem of the actual structure of interstellar ices.
APPENDIX A: ONLINE MATERIAL

In the online material file, we provide (i) a detailed discussion of the computational details (basis set, k-space sampling) and of the equation we used in this work; (ii) the energetic and vibrational properties of different CO/CO and CO/H₂O trimer complexes; (iii) the resume of all computed IR vibrational wavenumbers and intensities for all models we have developed, as well as a symmetry analysis for sIII.grp1 and sIII.grp2 CO-containing clathrates; and (iv) the neighbourhood analysis for cage models hosting only one CO molecule per cage.

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