Stability of spherically symmetric solutions in modified theories of gravity

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In recent years, a number of alternative theories of gravity have been proposed as possible resolutions of certain cosmological problems or as toy models for possible but heretofore unobserved effects. However, the implications of such theories for the stability structures such as stars have not been fully investigated. We use our “generalized variational principle”, described in a previous work [1], to analyze the stability of static spherically symmetric solutions to spherically symmetric perturbations in three such alternative theories: Carroll et al.’s $f(R)$ gravity, Jacobson & Mattingly’s “Einstein-æther theory”, and Bekenstein’s TeVeS. We find that in the presence of matter, $f(R)$ gravity is highly unstable; that the stability conditions for spherically symmetric curved vacuum Einstein-æther backgrounds are the same as those for linearized stability about flat spacetime, with one exceptional case; and that the “kinetic terms” of vacuum TeVeS are indefinite in a curved background, leading to an instability.

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I. INTRODUCTION

The idea of using a variational principle to bound the spectrum of an operator is familiar to anyone who has taken an undergraduate course in quantum mechanics: given a Hamiltonian $H$ whose spectrum is bounded below by $E_0$, we must have for any normalized state $|\psi\rangle$ in our Hilbert space $E_0 \leq \langle \psi | H | \psi \rangle$. We can thus obtain an upper bound on the ground state energy of the system by plugging in various “test functions” $|\psi\rangle$, allowing these to vary, and finding out how low we can make the expectation value of $H$.

This technique is not particular to quantum mechanics; rather, it is a statement about the properties of self-adjoint operators on a Hilbert space. In particular, we can make a similar statement in the context of linear field theories. Suppose we have a second-order linear field theory, dependent on some background fields and some dynamical fields $\psi^\alpha$, whose equations of motion can be put into the form

\[-\frac{\partial^2}{\partial t^2} \psi^\alpha = \mathcal{T}^{\alpha\beta} \psi^\beta\]  

(1)

for some linear spatial differential operator $\mathcal{T}$. Suppose, further, that we can find an inner product $(\cdot,\cdot)$ such that $\mathcal{T}$ is self-adjoint under this inner product, i.e.,

\[(\psi_1, \mathcal{T} \psi_2) = (\mathcal{T} \psi_1, \psi_2)\]  

(2)

for all $\psi_1$ and $\psi_2$. Then the spectrum of $\mathcal{T}$ will correspond to the squares of the frequencies of oscillation of the system; and we can obtain information about the lower bound of this spectrum via a variational principle of the form

\[\omega^2_{\beta} \leq \frac{(\psi, \mathcal{T} \psi)}{(\psi, \psi)}.\]  

(3)

Moreover, if we find that we can make this quantity negative for some test function $\psi$, then the system is unstable [2]; and we can obtain an upper bound on the timescale of this instability by finding how negative $\omega^2_{\beta}$ must be.

One might think that this would be a simple way to analyze any linear (or linearized) field theory; in practice, however, it is easier said than done. A serious problem arises when we consider linearizations of covariant field theories. In such theories, the linearized equations of motion are not all of the form (1); instead, we obtain non-deterministic equations (due to gauge freedom) and equations which are not evolution equations but instead relate the initial data for the fields $\psi^\alpha$ (since covariant field theories on a spacetime become constrained when decomposed into “space + time”.) Further, even if we have equations of the form (1), it is not always evident how to find an inner product under which the time-evolution operator $\mathcal{T}$ is self-adjoint (or even if such an inner product exists.) Thus, it would seem that this method is of limited utility in the context of linearized field theories.

In spite of these difficulties, Chandrasekhar successfully derived a variational principle to analyse the stability of spherically symmetric solutions of Einstein gravity with perfect-fluid sources [3, 4]. The methods used in these works to eliminate the constraint equations, put the equations in the form (1), and obtain an inner product seemed rather particular to the theory he was examining, and it was far from certain that they would generalize to an arbitrary field theory. In a recent paper [1], we showed that these methods did in a certain sense “have to work out”, by describing a straightforward procedure by which the gauge could be fixed, the constraints could be solved, and an inner product could be obtained under which the resulting time-evolution operator $\mathcal{T}$ is self-adjoint. We review this procedure in Section II of the paper. We then use this “generalized variational principle” to analyze three alternative theories of gravity which have garnered some attention in recent years: $f(R)$ gravity, the current interest in which was primarily inspired
by the work of Carroll et al.

II. REVIEW OF THE GENERALIZED VARIATIONAL PRINCIPLE

A. Symplectic dynamics

We first introduce some necessary concepts and notation. Consider a covariant field theory with an action of the form

\[ S = \int \mathcal{L} = \int \mathcal{L}[\Psi] \epsilon \] (4)

where \( \mathcal{L}[\Psi] \) is a scalar depending on some set \( \Psi \) of dynamical tensor fields including the spacetime metric. (For convenience, we will describe the gravitational degrees of freedom using the inverse metric \( g^{ab} \) rather than the metric \( g_{ab} \) itself.) To obtain the equations of motion for the dynamical fields, we take the variation of the four-form \( \mathcal{L} \epsilon \) with respect to the dynamical fields \( \Psi \):

\[ \delta (\mathcal{L} \epsilon) = (\mathcal{E}_\Psi \delta \Psi + \nabla_a \theta^a [\Psi, \delta \Psi]) \epsilon \] (5)

where a sum over all fields comprising \( \Psi \) is implicit in the first term. Requiring that \( \delta S = 0 \) under this variation then implies that the quantities \( \mathcal{E}_\Psi \) vanish for each dynamical field \( \Psi \).

The second term in (5) defines the vector field \( \theta^a [\Psi, \delta \Psi] \). The three-form \( \theta \) dual to this vector field (i.e., \( \theta_{abcd} = \theta^a \epsilon_{abcd} \)) is the “symplectic potential current.” Taking the antisymmetrized second variation of this quantity, we then obtain the symplectic current three-form \( \omega \) for the theory:

\[ \omega[\Psi; \delta_1 \Psi, \delta_2 \Psi] = \delta_1 \theta[\Psi, \delta_2 \Psi] - \delta_2 \theta[\Psi, \delta_1 \Psi]. \] (6)

In terms of the vector field \( \omega^a \) dual to \( \omega \), this can also be written as

\[ \omega^a \epsilon_{abcd} = \delta_1 (\theta^a_2 \epsilon_{abcd}) - \delta_2 (\theta^a_1 \epsilon_{abcd}). \] (7)

The symplectic form for the theory is then obtained by integrating the pullback of this three-form over a space-like three-surface \( \Sigma \):

\[ \Omega[\Psi; \delta_1 \Psi, \delta_2 \Psi] = \int_{\Sigma} \omega[\Psi; \delta_1 \Psi, \delta_2 \Psi]. \] (8)

If we define \( n^a \) as the future-directed timelike normal to \( \Sigma \) and \( e \) to be the induced volume three-form on \( \Sigma \) (i.e., \( \epsilon_{abcd} = n^a \epsilon_{abcd} \)), this can be written in terms of \( \omega^a \) instead:

\[ \Omega = - \int_{\Sigma} (\omega^a n_a) e. \] (9)

In performing the calculations which follow, it is this second expression for \( \Omega \) which will be most useful to us.

B. Obtaining a variational principle

In [1], we presented a procedure by which a variational principle for spherically symmetric perturbations of static, spherically symmetric spacetimes could generally be obtained. For our purposes, we will outline this method; further details can be found in the original paper.

The method described in [1] consists of the following steps:

1. Vary the action to obtain the equations of motion \( (\mathcal{E}_G)_{ab} \) corresponding to the variation of the metric, as well as any other equations of motion \( \mathcal{E}_A \) corresponding to the variations of any matter fields present. This variation will also yield the dual \( \theta^a \) of the symplectic current potential; take the antisymmetrized variation of this quantity (as in (7)) to obtain the symplectic form (10).

2. Fix the gauge for the metric, and choose an appropriate set of spacetime functions to describe the matter fields. Throughout this paper, we will choose our coordinates such that the metric takes the form

\[ ds^2 = -e^{2\Phi(r,t)} dt^2 + e^{2\Lambda(r,t)} dr^2 + r^2 d\Omega^2 \] (10)

for some functions \( \Phi \) and \( \Lambda \). Our spacetimes will be static at zero-order (i.e., \((\partial/\partial t)^a \) is a Killing vector field at zero-order), but non-static in first-order perturbations. We will therefore have \( \Phi(r,t) = \Phi(r) + \phi(r,t) \), where \( \phi \) is a first-order quantity; similarly, we define the first-order quantity \( \lambda \) such that \( \Lambda(r,t) = \Lambda(r) + \lambda(r,t) \). In other words,

\[ \delta g_{tt} = 2e^{-2\Phi(r)} \phi(r,t) \] (11)

and

\[ \delta g_{rr} = -2e^{-2\Lambda(r)} \lambda(r,t). \] (12)

Similarly, all matter fields will be static in the background, but possibly time-dependent at first order.

\[ 1 \text{ Such a set of coordinates can always be found for a spherically symmetric spacetime [10].} \]
3. Write the linearized equations of motion and the symplectic form in terms of these perturbational fields (metric and matter.)

4. Solve the linearized constraints. One of the main results of \[1\] was to show that this can be done quite generally for spherically symmetric perturbations off of a static, spherically symmetric background. Specifically, suppose the field content \(\Psi\) of our theory consists of the inverse metric \(g^{\mu\nu}\) and a single tensor field \(A^{a_1 \ldots a_n b_1 \ldots b_m}\). (The generalization to multiple tensor fields is straightforward.) Let \((\mathcal{E}_A)_{a_1 \ldots a_n b_1 \ldots b_m}\) denote the equation of motion associated with \(A^{a_1 \ldots a_n b_1 \ldots b_m}\). We define the constraint tensor as

\[
C_{cd} = 2(\mathcal{E}_A)_{cd} - g_{ce} \sum_i A^{a_1 \ldots a_n b_1 \ldots b_m} \mathcal{E}_{a_1 \ldots a_n b_1 \ldots b_m} + g_{ce} \sum_i A^{a_1 \ldots c \ldots a_n b_1 \ldots b_m} \mathcal{E}_{a_1 \ldots c \ldots a_n b_1 \ldots b_m}
\]  

(13)

where the summations run over all possible index “slots”, from 1 to \(n\) and from 1 to \(m\) for the first and second summation respectively. It can then be shown that if the background equations of motion hold, and the matter equation of motion also holds to first order, the perturbations of the tensor \(C_{ab}\) will satisfy

\[
\frac{\partial F}{\partial t} = -r^2 e^{-\lambda} \delta C_{tt}
\]  

(14)

and

\[
\frac{\partial F}{\partial r} = -r^2 e^{\lambda} \delta C_{tt}
\]  

(15)

for some quantity \(F\) which is linear in the first-order fields. Moreover, the first-order constraint equations \(\delta C_{tt} = \delta C_{rr} = 0\) will be satisfied if and only if \(F = 0\). We can then solve this equation algebraically for one of our perturbational fields, usually the metric perturbation \(\lambda\). We will refer to the equation \(F = 0\) as the preconstraint equation.

5. Eliminate the metric perturbation \(\phi\) from the equations. As \(\phi\) cannot appear without a radial derivative (due to residual gauge freedom)\(^2\), we must find an algebraic equation for \(\partial \phi / \partial r\). The first-order equation \(\delta C_{tt} = 0\), will, in general, serve this purpose \[1\]. Use the above relations for \(\lambda\) and \(\partial \phi / \partial r\) to eliminate the metric perturbations completely from the perturbational equations of motion and the symplectic form, leaving a “reduced” set of equations of motion and a “reduced” symplectic form solely in terms of the matter variables.

6. Determine if the reduced equations of motion take the form \[1\]. If so, read off the time-evolution operator \(T\).

7. Determine if the symplectic form, written in terms of the reduced dynamical variables \(\psi^\alpha\), is of the form

\[
\Omega(\Psi; \psi^\alpha_1, \psi^\alpha_2) = \int_{\Sigma} W_{\alpha \beta} \left( \frac{\partial \psi^\alpha_1}{\partial t} \psi^\beta_2 - \frac{\partial \psi^\alpha_2}{\partial t} \psi^\beta_1 \right)
\]

for some three-form \(W_{\alpha \beta}\). Then (as we showed in \[1\]) we must have \(W_{\alpha \beta} = W_{\beta \alpha}\) and

\[
\int_{\Sigma} W_{\alpha \beta} \psi^\alpha_1 T^\beta \gamma_2 = \int_{\Sigma} W_{\alpha \beta} \psi^\alpha_2 T^\beta \gamma_1.
\]

(17)

8. Determine whether \(W_{\alpha \beta}\) is positive definite, in the sense that

\[
\int_{\Sigma} W_{\alpha \beta} \psi^\alpha \psi^\beta \geq 0
\]

for all \(\psi^\alpha\), with equality holding only when \(\psi^\alpha = 0\).\(^3\) If this is the case, we can define an inner product on the space of all reduced fields:

\[
(\psi_1, \psi_2) \equiv \int_{\Sigma} W_{\alpha \beta} \psi^\alpha_1 \psi^\beta_2.
\]

(19)

Equation (17) then shows that \(T\) is a symmetric operator under this inner product. Thus, we can write down our variational principle of the form \[3\].

It is important to note that each step of this procedure is clearly delineated. While the procedure can fail at certain steps (the reduced equations of motion can fail to be of the form \[1\], for example, or \(W_{\alpha \beta}\) can fail to be positive definite), there is not any “art” required to apply this procedure to an arbitrary covariant field theory. (In practice, as we shall see, there are certain shortcuts that may arise which we can exploit; however, the “long way” we have described here will still work.) In the following three sections, we will use this formalism to analyze the stability of \(f(R)\) gravity, Einstein-êther theory, and TeVeS.

\(^2\) The exception to this statement is when the matter fields under consideration have non-zero components in the \(t\)-direction. In this case, \(\phi\) and its derivatives can appear in linear combination with perturbations of the matter fields; however, a redefinition of the matter fields will suffice to eliminate such cases.

\(^3\) If \(W_{\alpha \beta}\) is negative definite in this sense, we use the negative of this quantity as our inner product, and the construction proceeds identically.
III. \( f(R) \) Gravity

A. Theory

In \( f(R) \) gravity, the Ricci scalar \( R \) in the Einstein-Hilbert action is replaced by an arbitrary function of \( R \), leaving the rest of the action unchanged; in other words, the Lagrangian four-form \( \mathcal{L} \) is of the form

\[
\mathcal{L} = \frac{1}{16\pi} f(R) \epsilon + \mathcal{L}_{\text{mat}}[A, g^{ab}] 
\]

where \( A \) denotes the collection of matter fields, with tensor indices suppressed. Taking the variation of this action with respect to the metric, we obtain the equation of motion

\[
f'(R)R_{ab} - \frac{1}{2} f(R) g_{ab} - \nabla_a \nabla_b f'(R) + g_{ab} \Box f'(R) = 8\pi T_{ab}
\]

where \( T_{ab} \), given by

\[
\delta \mathcal{L}_{\text{mat}} = - \frac{1}{2} (T_{ab} \delta g^{ab}) \epsilon,
\]

is the matter stress-energy tensor, and \( f'(R) = df/dR \).

This equation is fourth-order in the metric, and as such is somewhat difficult to deal with. We can reduce this fourth-order equation to two second-order equations using an equivalent scalar-tensor theory [11, 12].\(^4\) This equivalent theory contains two dynamical gravitational variables, the inverse metric \( g^{ab} \) and a scalar field \( \alpha \), in addition to the matter fields. The Lagrangian is given by

\[
\mathcal{L} = \frac{1}{16\pi} \left( f'(\alpha) R + f(\alpha) - \alpha f'(\alpha) \right) \epsilon + \mathcal{L}_{\text{mat}}[A, g^{ab}].
\]

Varying the gravitational part of action with respect to \( g^{ab} \) and \( \alpha \) gives us

\[
\delta \mathcal{L} = (\mathcal{E}_G)_{ab} \delta g^{ab} + \mathcal{E}_\alpha \delta \alpha + \mathcal{E}_A \delta A + \nabla_a \theta^a \epsilon
\]

where \( \mathcal{E}_A \) denotes the matter equations of motion,

\[
(\mathcal{E}_G)_{ab} = \frac{1}{16\pi} \left[ f'(\alpha) G_{ab} - \nabla_a \nabla_b f'(\alpha) + g_{ab} \Box f'(\alpha) - \frac{1}{2} g_{ab} (f(\alpha) - \alpha f'(\alpha)) - 8\pi T_{ab} \right]
\]

and

\[
\mathcal{E}_\alpha = f''(\alpha)(R - \alpha),
\]

and the vector \( \theta^a \) is our symplectic potential current:

\[
\theta^a = f'(\alpha) \theta^a_{\text{Ein}} + \theta^a_{\text{mat}} + \frac{1}{16\pi} \left( \frac{1}{2} (\nabla_b f'(\alpha)) \delta g^{ab} - (\nabla^a f'(\alpha)) g_{bc} \delta g^{bc} \right).
\]

The vector \( \theta^a_{\text{mat}} \) above is the symplectic potential current resulting from variation of \( \mathcal{L}_{\text{mat}} \), and \( \theta^a_{\text{Ein}} \) is the symplectic potential current for pure Einstein gravity, i.e.,

\[
\theta^a_{\text{Ein}} = \frac{1}{16\pi} \left( g_{bc} \nabla^a \delta g^{bc} - \nabla_b \delta g^{ab} \right).
\]

The equations of motion are then given by \( (\mathcal{E}_G)_{ab} = 0 \) and \( \mathcal{E}_\alpha = 0 \). Assuming that \( f''(\alpha) \neq 0 \), this second equation implies that \( \alpha = 0 \), and substituting this relation into (23) yields the equation of motion obtained in [21]. Hereafter, we will use the form of the equations obtained from the action (23).

B. Obtaining a variational principle

Before we examine the stability of spherically symmetric static solutions in \( f(R) \) gravity with perfect fluid matter, we must first consider the question of whether physically realistic solutions exist. In particular, do there exist solutions to the equations of motion which reproduce Newtonian gravity, up to small relativistic corrections? There has recently been a good deal of debate on this subject [14, 15, 16, 17, 18]. We will not comment directly on this controversy here except to say that if \( f(R) \) gravity (or any theory) does not allow interior solutions with \( R \approx 8\pi G \rho \) to be matched to exterior solutions with \( R \) close to zero, then it is difficult to see how such a theory could reproduce Newtonian dynamics in a nearly-flat spacetime. We will therefore give the theory the “benefit of the doubt,” and assume that such solutions exist.

An initial attempt to address the stability of such solutions was made by Dolgov and Kawasaki [19]; their perturbation analysis implied that stars would be extremely unstable in Carroll et al.’s \( f(R) \) gravity, with a characteristic time scale of approximately 10^{-16} seconds. Their results, while suggestive, nevertheless failed to take into account the constrained nature of the theory: the stress-energy tensor, the metric, and the scalar \( \alpha \) cannot all be varied independently. The imposition of a constraint can, of course, change whether or not a given system is stable; as a trivial example, consider a particle moving in the potential \( V(x, y) = x^2 - y^2 \), with and without the constraint \( y = \text{const.} \).\(^5\) In what follows, we will show that Dolgov and Kawasaki’s conclusion is, nonetheless,\(^6\)

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\(^4\) This equivalence was also used in the special case \( f(R) = R - 2\Lambda + \alpha R^2 \) in [13].

\(^5\) In certain scalar-tensor theories [20], the scalar perturbations decouple from the metric and matter perturbations; it is then legitimate in such theories to “ignore” the constraints if we concern ourselves only with the scalar field. However, this decoupling does not occur in \( f(R) \) gravity theory. To see this,
The Lagrangian is then given by

\[ \mathcal{L} = -\varrho(\nu)\varepsilon \]  

(30)

where \( \nu \), the “number density” of the fluid, is given by

\[ \nu^2 = \frac{1}{6} N_{abc} N^{abc}. \]  

(31)

In turn, \( N_{abc} \), the “number current” of the fluid, is given by

\[ N_{abc} = N_{ABC}(X) \nabla_a X^A \nabla_b X^B \nabla_c X^C. \]  

(32)

We will be purely concerned with spherically symmetric solutions and radial perturbations of the fluid; thus, we will take our Lagrangian coordinates to be of the form

\[ X^R = r, \quad X^\theta = \theta, \quad X^\Phi = \varphi \]  

(33)

in the background, and consider only perturbations \( \delta X^R = \delta X^\Theta = \delta X^\Phi = 0 \).

We can then easily obtain the background equations of motion; these are

\[ (E_G)_{tt} = e^{2\Phi - 2\Lambda} \left[ -\frac{\partial^2}{\partial r^2} f'(\alpha) + \left( \frac{\partial \Lambda}{\partial r} - \frac{2}{r} \right) \frac{\partial}{\partial r} f'(\alpha) \right] \]

\[ + \left( \frac{2 \partial \Lambda}{r} + \frac{1}{r^2} (e^{2\Lambda} - 1) \right) f'(\alpha) \]

\[ - \frac{1}{2} e^{2\Phi} \left( \alpha f'(\alpha) - f(\alpha) \right) - e^{2\Phi} \varrho = 0 \]  

(34a)

\[ (E_G)_{\theta\theta} = r^2 e^{-2\Lambda} \left[ \frac{\partial^2}{\partial r^2} f'(\alpha) + \left( \frac{\partial \Lambda}{\partial r} - \frac{1}{r} \right) \frac{\partial}{\partial r} f'(\alpha) \right] \]

\[ + \left( \frac{2 \partial \Lambda}{r^2} + \frac{1}{r^2} (e^{2\Lambda} - 1) \right) f'(\alpha) \]

\[ - \frac{1}{2} (\alpha f'(\alpha) - f(\alpha)) - r^2 (\varrho' - \varrho) = 0 \]  

(34b)

\[ (E_G)_{\phi\phi} = e^{2\Phi} \left[ \frac{\partial^2}{\partial r^2} f'(\alpha) + \left( \frac{\partial \Lambda}{\partial r} - \frac{2}{r} \right) \frac{\partial}{\partial r} f'(\alpha) \right] \]

\[ + \left( \frac{2 \partial \Lambda}{r} + \frac{1}{r^2} (e^{2\Lambda} - 1) \right) f'(\alpha) \]

\[ - \frac{1}{2} e^{2\Phi} \left( \alpha f'(\alpha) - f(\alpha) \right) - e^{2\Phi} \varrho = 0 \]  

(34c)

\[ (E_X)_t = \frac{\varrho}{\varrho'} \frac{\partial \Phi}{\partial r} + \frac{\partial \Phi}{\partial r} \varrho' = 0 \]  

(34d)

where \( \varrho' = d\varrho/d\nu \). These equations are not all independent; in particular, the Bianchi identity implies that \( E_G \) is automatically satisfied if the other four equations are satisfied. Note that under the substitutions \( \nu = \rho, \varrho' = \varrho' \rightarrow \varrho \), the “matter terms” in these equations take on their familiar forms for a perfect fluid.

We next obtain the symplectic form for the theory.

The contribution \( \theta^a_{\text{grav}} \) to the symplectic potential current from the gravitational portion of the action is given by (27); to obtain the symplectic form, we then take the antisymmetrized second variation of \( \theta \), as in (7). Performing this variation, we find that

\[ \omega^a_{\text{grav}} = f'(\alpha) \omega^a_{\text{Ein}} + \frac{1}{2} \delta_1 g^{bc} \delta_2 g^{ad} g_{bc} \nabla_d (f'(\alpha)) \]

\[ + \left[ \delta_1 (f'(\alpha)) \nabla_b \delta_2 g^{cd} - \nabla_b (\delta_1 (f'(\alpha))) \delta_2 g^{cd} \right] \times (g^{ab} g_{cd} - \delta^{a} \delta^{bd} - \delta^{b} \delta^{ad}) \]  

(35)

where \( \omega^a_{\text{Ein}} \) is defined, analogously to \( \theta^a_{\text{Ein}} \), to be the symplectic current associated with pure Einstein gravity. This is equal to (21)

\[ \omega^a_{\text{Ein}} = S^{a}_{bc} d_{ef} (\delta_2 g^{bc} \nabla_d \delta_1 g^{ef} - \delta_1 g^{bc} \nabla_d \delta_2 g^{ef}), \]  

(36)

where

\[ S^{a}_{bc} d_{ef} = \frac{1}{16\pi} \left( \delta^a e \delta^c g_{bf} - \frac{1}{2} g^{ad} g_{bc} g_{ef} \right) - \frac{1}{2} \delta^a \delta^c g_{ef} g_{bc} + \frac{1}{2} g^{ad} g_{bc} g_{ef}). \]  

(37)
The contribution to the symplectic current from the matter terms in the Lagrangian coordinate formalism was calculated in \[1\]; we simply cite the result for the t-component of $\omega^t$ in a static background here:

$$t_a \omega^a_{\text{mat}} = -t_a \frac{g'}{2\nu} N_{ABC} \nabla_b X^B \nabla_c X^C \left[ \delta_j^a \delta^d_2 X^A N_{d b c} \right.$$  
$$+ 3 \delta_2 X^A \nabla^a (N_{D E F} \delta_1 X^D) \nabla^b X^K \nabla^c X^F - [1 \leftrightarrow 2] \right],$$

where we have defined the tensor indices only (not the fluid-space indices). Writing out $\omega^t$ in terms of our perturbational variables, we have

$$\omega^t = \nu' e^{2\Lambda - 2\Phi} \left( \frac{\partial \xi_1}{\partial r} \xi_2 - \frac{\partial \xi_2}{\partial r} \xi_1 \right)$$
$$+ e^{-2\Phi} \left( b_1 \frac{\partial \lambda_2}{\partial r} - \frac{\partial b_1}{\partial r} \lambda_2 - b_2 \frac{\partial \lambda_1}{\partial r} + \frac{\partial b_2}{\partial r} \lambda_1 \right),$$

where we have defined $b = \delta(f'(\alpha)) = f''(\alpha) \delta \alpha$. (Note that the t-component of the symplectic current about a spherically symmetric static solution vanishes in pure Einstein gravity, i.e., $\omega^t_E = 0$.)

The first step towards obtaining a variational principle is to solve the linearized constraints. To do this, we must calculate the constraint tensor $C_{ab}$, as defined in \[15\]. Since all the fields in our theory other than the metric are scalars with respect to the spacetime metric, the resulting expression is particularly simple:

$$C_{ab} = 2(\mathcal{E}_G)_{ab},$$

where $(\mathcal{E}_G)_{ab}$, the metric equation of motion, is given by \[26\]. We can then find algebraic equations for $\lambda$ and $\partial \phi / \partial r$. The algebraic equation for $\lambda$ will be given by the solution of the equation $F = 0$, where $F$ is given by \[14\] and \[15\]. Similarly, an algebraic equation for $\partial \phi / \partial r$ can be found by solving the equation

$$\delta C_{rr} = (\delta \mathcal{E}_G)_{rr} = 0.$$

The situation is thus very similar to the example of pure Einstein gravity coupled to perfect fluid matter described in \[1\]; only the precise form of the perturbational equations of motion $F = 0$ and $(\delta \mathcal{E}_G)_{rr} = 0$ are different. In the $f(R)$ gravity case, these equations are

$$F = S \lambda - \frac{\partial b}{\partial r} + \frac{\partial \phi}{\partial r} b - 8\pi e^{2\Lambda} \nu' \xi$$

and

$$(\delta \mathcal{E}_G)_{rr} = e^{-2\Lambda} S \frac{\partial \phi}{\partial r} - e^{-2\Phi} \frac{\partial^2 b}{\partial r^2} + e^{-2\Lambda} \left( \frac{\partial \phi}{\partial r} + 2 \frac{g'}{r} \right) \frac{\partial b}{\partial r}$$
$$- e^{-2\Lambda} \left( \frac{2}{r} \frac{\partial \phi}{\partial r} + \frac{\partial \lambda}{\partial r} \right) + \left( e^{-2\Lambda} \frac{2}{r^2} \frac{\partial \phi}{\partial r} - \frac{\partial \lambda}{\partial r} \right) \frac{1}{r^2} \left( 1 - e^{-2\Lambda} \right) + \frac{\alpha}{2} \right) b$$
$$- 8\pi \nu' \frac{\partial \xi}{\partial r} + \left( \frac{\partial \lambda}{\partial r} + \frac{2}{r} \frac{\partial \nu}{\partial r} \right) \xi - \lambda, \quad \text{(43)}$$

where we have introduced the new quantity

$$S = \frac{\partial}{\partial r} f'(\alpha) + 2 \frac{g'}{r} f'(\alpha)$$

for notational convenience. Equation \[42\] then implies that our algebraic equation for $\lambda$ is

$$\lambda = S^{-1} \left( \frac{\partial b}{\partial r} - \frac{\partial \phi}{\partial r} b + 8\pi e^{2\Lambda} \nu' \xi \right).$$

We could plug this result into \[43\] to obtain an algebraic equation for $\partial \phi / \partial r$; however, the resulting expression is somewhat complicated. In fact, we can derive a simpler expression for $\partial \phi / \partial r$. To do so, we note that $(\mathcal{E}_G)_{ab} = \frac{1}{2} g_{ab} (\mathcal{E}_G) c^c = 0$ if and only if $(\mathcal{E}_G)_{ab} = 0$. In the case of $f(R)$ gravity, this equation is given by

$$f'(\alpha) R_{ab} - \frac{1}{2} g_{ab} (f(\alpha) - f'(\alpha) \alpha) - \nabla_a \nabla_b f'(\alpha)$$
$$+ \frac{1}{2} g_{ab} (f(\alpha) + f'(\alpha) \alpha) - 8\pi \left( T_{ab} - \frac{1}{2} g_{ab} T^c c \right) = 0.$$

(46)

Using the trace of this equation to eliminate the $\Box f'(\alpha)$ term, along with the equation $R = \alpha$, we can show that

$$f'(\alpha) R_{ab} + \frac{1}{6} g_{ab} (f(\alpha) - 2f'(\alpha) \alpha) - \nabla_a \nabla_b f'(\alpha)$$
$$- \frac{1}{2} \alpha (f(\alpha) - 2f'(\alpha) \alpha) - 8\pi \left( T_{ab} - \frac{1}{3} g_{ab} T^c c \right) = 0.$$

(47)

To zero-order, the $\theta\theta$-component of this equation is

$$\frac{1}{r} \frac{f'(\alpha)}{r} \left( \frac{\partial \lambda}{\partial r} - \frac{\partial \phi}{\partial r} + \frac{\partial \nu}{\partial r} \right) + \frac{1}{6} e^{2\Lambda} \left( f(\alpha) - 2\alpha f'(\alpha) \right) - \frac{\partial \nu}{\partial r} \frac{f'(\alpha)}{r} - \frac{8\pi}{3} \nu' e^{2\Lambda} \alpha = 0,$$

(48)

and to first order, it is
\[
\frac{1}{r} f'(\alpha) \left( \frac{\partial \lambda}{\partial r} - \frac{\partial \phi}{\partial r} + \frac{2}{r} e^{2\Lambda} \lambda \right) + \left( \frac{1}{r} \left( \frac{\partial \Delta}{\partial r} - \frac{\partial \Phi}{\partial r} \right) + \frac{1}{r^2} \left( e^{2\Lambda} - 1 \right) - \frac{1}{6} e^{2\Lambda} \left( \frac{f'(\alpha)}{f''(\alpha)} + 2\alpha \right) \right) \frac{b}{r} - \frac{1}{r} \frac{\partial b}{\partial r} + \frac{1}{3} \lambda \left( f(\alpha) - 2\alpha f'(\alpha) - 16\pi \varrho \right) \lambda - \frac{8\pi}{3} \varrho \nu e^{2\Lambda} \left( \frac{\partial \xi}{\partial r} + \left( \frac{\partial \Delta}{\partial r} + \frac{2}{r} + \frac{1}{\nu} \frac{\partial \nu}{\partial r} \right) \xi - \lambda \right) = 0. \tag{49}
\]

This equation, with \( \lambda \) given by (45), can then be solved algebraically for \( \partial \phi/\partial r \) in terms of the matter variables and their derivatives.

Our next step is to obtain the reduced matter equations of motion, i.e., eliminate the metric degrees of freedom \( \lambda \) and \( \partial \phi/\partial r \) from the remaining first-order equations of motion. These remaining equations of motion are the “scalar” equation \( R - \alpha = 0 \), which at first order is

\[
- \frac{\partial^2 \phi}{\partial r^2} + \frac{\partial \Phi}{\partial r} + \frac{\partial \Delta}{\partial r} \frac{\partial \phi}{\partial r} - 2 \frac{\partial \phi}{\partial r} + \frac{2}{r} \left( \frac{\partial \lambda}{\partial r} - \frac{\partial \phi}{\partial r} \right) + \frac{2}{r^2} e^{2\Lambda} \lambda + e^{2\Lambda - 2\Phi} \frac{\partial^2 \lambda}{\partial t^2} = e^{2\Lambda} \left( \alpha \lambda + \frac{b}{2 f''(\alpha)} \right) = 0 \tag{50}
\]

and the matter equation of motion, which as in the case of pure Einstein gravity is

\[
\varrho' \left[ - e^{2\Lambda - 2\Phi} \frac{\partial^2 \xi}{\partial t^2} + \frac{\partial \phi}{\partial r} \right] + \left( \frac{\partial}{\partial r} + \frac{\partial \phi}{\partial r} \right) \left[ \varrho' \nu \left( \frac{\partial \xi}{\partial r} + \left( \frac{1}{\nu} \frac{\partial \nu}{\partial r} + \frac{2}{r} \right) \xi - \lambda \right) \right] = 0. \tag{51}
\]

We can then eliminate the gravitational equations of motion using (45) and (49), leaving equations solely in terms of the matter variables \( b \) and \( \xi \). Equivalently, we can write our “reduced” equations in terms of \( b \) and a new variable \( \zeta \), defined by

\[
\zeta \equiv \xi - S^{-1} b. \tag{52}
\]

(The utility of this new variable will become evident when we obtain the reduced symplectic form.) Performing these algebraic manipulations, the resulting reduced evolution equation for \( \zeta \) takes the form

\[
e^{-2\Phi} \varrho' \frac{\partial^2 \zeta}{\partial t^2} = A_1 \frac{\partial \zeta}{\partial r} + A_2 \frac{\partial^2 \zeta}{\partial r^2} + A_3 \zeta + A_4 \frac{\partial b}{\partial r} + A_5 b \tag{53}
\]

where the coefficients \( A_i \) are dependent on the background fields. Similarly, the evolution equation for \( b \) takes the form

\[
e^{-2\Phi} \frac{\partial^2 b}{\partial t^2} = B_1 \frac{\partial b}{\partial r} + B_2 \frac{\partial^2 b}{\partial r^2} + B_3 b + B_4 \frac{\partial \zeta}{\partial r} + B_5 \zeta. \tag{54}
\]

We see from the form of the above equations that \( f(R) \) gravity has cleared another hurdle required to have a valid variational principle: the reduced equations do in fact take the form (11), containing only second derivatives in time and only up to second radial derivatives.

We can also apply these constraint equations to the symplectic form. For the symplectic form, we only need to eliminate \( \lambda \) using the first-order constraint equation (45). Performing an integration by parts to eliminate the mixed derivatives that result, and applying the background equations of motion, we obtain our reduced symplectic form:

\[
\Omega = 4\pi \int dr \, e^{\Lambda + \Phi} \left[ e^{-2\Phi} S^{-2} \frac{6}{r^2} f'(\alpha) \frac{\partial b_1}{\partial t} b_2 + e^{-2\Phi} \varrho' \nu \frac{\partial \zeta_1}{\partial t} \xi \right] - [1 \leftrightarrow 2]. \tag{55}
\]

As noted above (16), this reduced symplectic form defines a three-form \( W_{\alpha\beta} \). For a valid variational principle to exist, this \( W_{\alpha\beta} \) must be positive definite in the sense of (18). We can see that the \( W_{\alpha\beta} \) defined by (55) is positive definite if and only if \( f'(\alpha) > 0 \) and \( \varrho' \nu > 0 \) in our background solutions. The latter condition will hold for any matter satisfying the null energy condition, since \( \varrho' \nu = \rho + P \); however, the former condition must be checked in order to determine whether a valid variational principle exists.\(^6\) For the particular \( f(R) \) chosen by Carroll et al. (5), we have \( f'(R) = 1 + \mu R^2 > 0 \), and so \( W_{\alpha\beta} \) is always positive definite in the required sense for this choice of \( f(R) \).

All that remains is to actually write down the variational principle for \( f(R) \) gravity. As noted above (19), the variational principle will take the form

\[
\omega_0^2 \leq \frac{(\psi, T\psi)}{(\psi, \psi)} \tag{56}
\]

where \( T \) is the time-evolution operator. In our case, \( \psi \) denotes a two-component vector, whose components are the functions \( \zeta \) and \( b \). The inner product in which \( T \) is self-adjoint can be read off from (55), and thus the

\(^6\) Should \( f'(\alpha) \) fail to be positive in the background solutions, all is not necessarily lost; we can still attempt to analyze the reduced equations that we have obtained. See Section V for an example of such an analysis.
denominator of (56) will be

$$ (\psi, \psi) = 4\pi \int dr r^2 e^{1+\Phi} \left[ S^{-2} f'(\alpha) \frac{6 g''}{r^2} b^2 + e^{2\Lambda} g' \nu c^2 \right].$$

(57)

The numerator of (56) is, of course, rather more complicated. After multiple integrations by parts and applications of the background equations of motion, we can put this quantity in the form

$$ (\psi, T\psi) = 4\pi \int dr \left[ C_1 \left( \frac{db}{dr} \right)^2 + C_2 \left( \frac{d\zeta}{dr} \right)^2 + C_3 \left( \frac{ \zeta \partial_b}{r} - \left( \frac{d\zeta}{dr} \right) \right)^2 + C_4 b^2 + C_5 \zeta^2 + 2C_6 b\zeta \right],$$

(58)

where the $C_i$ coefficients are functions of the background fields. The exact forms of the $C_i$ are given in Appendix A.

C. Discussion

All of our results thus far have been independent of the choice of $f(R)$ (assuming, of course, that $f''(R) \neq 0$.) In the case of Carroll et al.'s $f(R)$ gravity, we can now use this variational principle to show that this theory is highly unstable for Newtonian solutions. For the choice $f(R) = R - \mu^2/R$, we have $f'(R) = 1 + \mu^4/R^2$ and $f''(R) = -2\mu^4/R^3$. Moreover, for a quasi-Newtonian stellar interior, we will have $\alpha = R \approx \rho$, where $\rho$ is the matter density in the star. In particular, this implies that

$$ C_4 \approx 2e^{\Phi+\Lambda} S^{-2} \frac{(f'(\alpha))^2}{f''(\alpha)} \approx -e^{\Phi+\Lambda} S^{-2} \frac{\rho^3}{\mu^4}. $$

(59)

since the above term will dominate over all the others in $C_4$. (Note that for a star with the density of the Sun, and the choice of $\mu \approx 10^{-27}$ m$^{-1}$ made by Carroll et al., $\rho/\mu^2 \approx 10^{10}$.) This implies that for a test function $\psi$ with $\zeta$ set to zero, we will have

$$ \omega_0^2 \lesssim \int dr e^{\Lambda+\Phi} S^{-2} \left[ -\frac{g''}{r^2} b^2 + 6 e^{-2\Lambda} \frac{(db)}{r(\bar{r})^2} \right],$$

(60)

(note that $f'(\alpha) = 1 + \mu^4/\rho^2 \approx 1$ in the stellar interior.) As a representative mass distribution, we take the Newtonian mass profile of an $n = 1$ polytrope:

$$ \rho(r) = \rho_0 \frac{R \sin\left( \frac{r}{R} \right)}{r},$$

(61)

where $R$ is the radius of the star and $\rho_0$ is its central density. We take $\rho_0$ to be of a typical stellar density, $\rho_0 \approx 10^{-24}$ metres$^{-3}$. Numerically integrating (59) with a test function of the form

$$ b = \begin{cases} 1 - \frac{r}{R} & r \leq R \\ 0 & r > R \end{cases},$$

(62)

we find that

$$ \omega_0^2 \lesssim -9 \times 10^{35} \text{metres}^2, $$

(63)

which corresponds to an instability timescale of $\tau \approx 4 \times 10^{27}$ seconds. This timescale is of the same magnitude as the instability found in (13).

We see, then, that for the $f(R)$ originally chosen by Carroll et al., the theory is extremely unstable in the presence of matter. Moreover, a similar argument will obtain for any choice of $f(R)$ for which quasi-Newtonian solutions exist and for which $f''(R)$ is sufficiently small and negative at stellar-density curvature scales. We can always pick a test function $b(r)$ lying entirely inside the stellar interior such that $\partial b/\partial r$ is of order $b/R$. Thus, if $|f''(\rho)/f'''(\rho)| \approx 1/R^2$ for a typical stellar density $\rho$, the $b^2$ term in (59) will dominate the $\partial b/\partial r^2$ term, and the resulting lower bound on $\omega_0$ will then be of order $f'(\alpha)/f''(\alpha)$. If the choice of $f(\alpha)$ results in this quantity being negative, quasi-Newtonian stellar solutions will be unstable in the corresponding theory. We can thus rule out any theory (e.g., $[5, 22]$) for which $f''(\rho)$ is sufficiently small and negative.

This result seems to be closely related, if not identical to, the “high-curvature” instabilities found in cosmological solutions by Song, Hu, and Sawicki [23]. Particularly telling is the fact that the instability timescale found in their work is proportional to $\sqrt{|f''(R)/f'(R)|}$, the same timescale found in the present work.

IV. EINSTEIN-ÆTER THEORY

A. Theory

Einstein-æther theory $[6, 7]$ was first formulated as a toy model of a gravitational theory in which Lorentz symmetry is dynamically broken. This theory contains, along with the metric $g_{ab}$, a vector field $u^a$ which is constrained (via a Lagrange multiplier $Q$) to be unit and timelike. The Lagrangian four-form for this theory is

$$ \mathcal{L} = \left( \frac{1}{16\pi} R + K_{cd} u^a \nabla_a u^c \nabla_b u^d + Q(u^a u_a + 1) \right) \epsilon $$

$$ + \mathcal{L}_{\text{mat}} [A, g^{ab}, u^a], $$

(64)

where $A$ denotes any matter fields present in the theory, and

$$ K_{cd} = c_1 g^{ab} g_{cd} + c_2 \delta^a e^{\delta^b} e^{\delta^d} + c_3 \delta^a g^{bd} - c_4 u^a u^b g_{cd}. $$

(65)

The $c_i$ constants determine the strength of the vector field’s coupling to gravity, as well as its dynamics.\footnote{Note that our definitions of the coefficients $c_i$ differ from those in $[5, 24, 25, 22]$; do as the respective metric signature conventions. The net result is that to translate between our results and those of the above papers, one must flip the signs of all four coefficients.} (Note
that in the case $c_3 = -c_1 > 0$ and $c_2 = c_4 = 0$, we have the conventional kinetic term for a Maxwell field; this special case was examined in [26], prior to Jacobson and Mattingly’s work.) In the present work, we will work in the “vacuum theory”, i.e., in the absence of matter fields $A$. Performing the variation of the Lagrangian four-form, we find that

$$\delta L = (E_G)_{ab} \delta g^{ab} + (E_u)_a \delta u^a + E_Q \delta Q + \nabla_a \theta^a \epsilon$$

(66)

where

$$(E_G)_{ab} = \frac{1}{16\pi} g_{ab} + \nabla_c \left( J_c(a u_b) + J_{ab} u^c - J_c(a u^b) \right)$$

$$+ c_1 \left( \nabla_a u^c \nabla_b u_c - \nabla_c u_a \nabla_b u^c \right) + c_4 \dot{u}_a \dot{u}_b$$

$$- Q u_a u_b - \frac{1}{2} g_{ab} J^c d \nabla_c u^d,$$  

(67a)

$$(E_u)_a = -2 \nabla_b J^b a - 2 c_4 u_b \nabla_a u^b + 2 Q u_a,$$  

(67b)

$$(E_Q) = u^a u_a + 1,$$  

(67c)

and

$$\theta^a = \theta_{\text{Ein}}^a + 2 J^a b \dot{u}^b + (J^a_{bc} u_c - J^a_{bc} u^c - J_{bc} u^a) \delta g^{bc}.$$  

(68)

In the above, we have introduced the notation $\dot{u}^a = u^b \nabla_b u^a$ and $J^a_{bc} = K^{ac} b \nabla_c u^d$; $\theta_{\text{Ein}}^a$ is defined by (28), as above. Equation (67c) is the constraint that $u^a$ is unit and timelike, while (67a) and (67b) are the equations of motion for $g^{ab}$ and $u^a$, respectively. If desired, we can eliminate the Lagrange multiplier $Q$ from these equations by contracting (67b) with $u^a$, resulting in the equation

$$Q = -u^a \nabla_b J^b a - c_4 u^a \dot{u}_a \nabla_c u^b.$$  

(69)

We now take the variation of the symplectic potential current to obtain the symplectic current. The resulting expression can be written in the form

$$\omega^a = \omega_{\text{Ein}}^a + \omega_{\text{vec}}^a$$  

(70)

where $\omega_{\text{Ein}}^a$ is the usual symplectic current for pure Einstein gravity (given by (33)), and $\omega_{\text{vec}}^a$ is obtained by taking the antisymmetrized variation of the last two terms in (68). The exact form of this tensor expression is rather complicated, and can be found in Appendix B.

### B. Obtaining a variational principle

As before, we are primarily concerned with the $t$-component of $\omega^a$, and its form in terms of the perturbational variables. We will use the usual spherical gauge (11) for $g^{ab}$. We will further assume that $u^a \propto \ell^a$ in the background solution (the so-called “static æther” assumption), and that to first order we have

$$u^a = (e^{-\Phi} - \phi e^{-\Phi}, \nu, 0, 0),$$

(71)

i.e., the first-order perturbation to $u^t$ is $-\phi e^{-\Phi}$ while the first-order perturbation to $u^r$ is $\nu$. (Note that under these conventions, $g_{ab} u^a u^b = -1$ to first order, as required.) The background equations of motion can then be calculated to be

$$(E_G)_{tt} = e^{2\Phi - 2\Lambda} \left[ \frac{2 \partial \Delta}{\partial r} + \frac{1}{r^2} \left( e^{2\Lambda} - 1 \right) \right]$$

$$+ c_{14} \left( \frac{\partial^2 \Phi}{\partial r^2} + \frac{1}{2} \left( \frac{\partial \Phi}{\partial r} \right)^2 - \frac{\partial \Phi}{\partial r} \left( \frac{\partial \Lambda}{\partial r} - \frac{2}{r} \right) \right),$$

(72)

and

$$(E_G)_{rr} = \frac{2 \partial \Phi}{\partial r} - \frac{1}{r^2} \left( e^{2\Lambda} - 1 \right) - \frac{c_{14}}{2} \left( \frac{\partial \Phi}{\partial r} \right)^2,$$  

(73)

where we have defined $c_{14} = c_1 + c_4$ and used the equation

$$Q = e^{-2\Lambda} \left( c_3 - 2 c_4 \left( \frac{\partial \Phi}{\partial r} \right)^2 - c_3 \frac{\partial^2 \Phi}{\partial r^2} \right)$$

(74)

to eliminate $Q$. The final equation of motion, $(E_u)_a = 0$, is satisfied trivially if the æther is static.

In terms of these variables, the $t$-component of $\omega^a$ can be calculated to be

$$\omega^t = 2 e^{-2\Phi} \left[ c_{123} \frac{\partial \lambda_1}{\partial t} \lambda_2 + c_{123} e^\Phi \frac{\partial \lambda_1}{\partial \nu} \lambda_2 \right]$$

$$+ e^\Phi \left( c_{123} \frac{\partial \Lambda}{\partial t} + \left( c_{123} - c_{14} \right) \frac{\partial \Phi}{\partial r} + c_2 \frac{2}{r} \right) v_1 \lambda_2$$

$$+ c_{14} e^\Phi v_1 \frac{\partial \phi_2}{\partial t} + e^{2\Lambda} c_{14} \frac{\partial \nu_2}{\partial t} \right] - \left[ 1 \leftrightarrow 2 \right]$$

(75)

where we have defined $c_{123} = c_1 + c_2 + c_3$. In what follows, we will assume that $c_1 \neq 0$, and, except where otherwise noted, that $c_{123} \neq 0$ as well.

Our next step, as usual, is to solve the constraints. However, in this theory we have the added complication of the presence of a vector field. This means, in particular, that the tensor $C_{ab}$ (as defined in (13)) is not merely proportional to $(E_G)_{ab}$, as in the previous section, but is instead

$$C_{ab} = 2 (E_G)_{ab} + u_a (E_u)_b$$

(76)

$$= \frac{1}{8\pi} G_{ab} + 2 \nabla_c \left( J_{(ab)} u^c - J_{(c) b} u^c \right) + 2 \nabla_c J^c_{a b} + 2 c_1 \nabla_a u^c \nabla_b u^d - 2 c_4 (\nabla_a u^c \nabla_b u^d - \nabla_a \nabla_b u^d) - g_{ab} J^c d \nabla_c u^d.$$  

(77)

Writing out $\delta C_{rt}$ in terms of our perturbational quantities, we find that it is indeed a total time derivative,
with
\[ F = 2r^2 e^\Phi - \lambda \left( \frac{2}{r} - c_{14} \frac{\partial \Phi}{\partial r} \right) \lambda + c_{14} e^{2\lambda - \Phi} \frac{\partial \Phi}{\partial t} + c_{14} e^{-\Phi} \frac{\partial \Phi}{\partial r}. \] (78)

Note that this quantity depends on \( \phi; \) \( \phi \) and its derivatives can appear in the preconstraint equation \( F = 0 \) when our background tensor fields have non-vanishing \( t \)-components, as noted in Footnote 2.

The remaining non-trivial equations of motion are
\[ \frac{1}{2} \left( \delta E_a \right)_r = c_{14} e^{2\lambda - 2\Phi} \frac{\partial^2 \psi}{\partial t^2} - c_{123} e^{2\Phi} \frac{\partial^2 \psi}{\partial r^2} - c_{123} e^{-\Phi} \frac{\partial ^2 \lambda}{\partial r \partial t} + c_{14} e^{-\Phi} \frac{\partial ^2 \phi}{\partial r \partial t} - c_{123} \left( \frac{2}{r} + \frac{\partial \lambda}{\partial r} + \frac{\partial \phi}{\partial r} \right) \frac{\partial \psi}{\partial r} + e^{-\Phi} \left( c_{123} (c_{123} - c_{14}) \frac{\partial \phi}{\partial r} - (c_1 + c_3) \frac{2}{r} \frac{\partial \lambda}{\partial r} \right) c_{123}^2 \left( \frac{\partial^2 \Phi}{\partial t^2} - \frac{\partial^2 \lambda}{\partial t \partial r} \right) c_{123} + \frac{2}{r} \frac{(c_1 + c_3) \frac{\partial \Lambda}{\partial r} + (c_3 - c_1) \frac{\partial \Phi}{\partial r}}{c_{123}^2} \frac{\partial \psi}{\partial r}. \] (81)

And
\[ \frac{\partial \psi}{\partial t} - e^{2\Phi - 2\lambda} \frac{2}{r^2} \left( \frac{2}{c_14} + 1 \right) \lambda = 0 \] (82)

where we have defined the constant \( a \) to be
\[ a = (c_1 + c_3 + 1)(c_{123} + 2c_2 - 2). \] (85)

We can then use this equation and (81) to write down a single second-order time-evolution equation for \( \psi \):
\[ e^{2\lambda - 2\Phi} \frac{r^2}{2 + c_{14}} \frac{\partial^2 \psi}{\partial t^2} + \frac{r^2}{a} \left[ \frac{c_{123}}{2} \left( \frac{\partial^2 \psi}{\partial r^2} + \frac{\partial \psi}{\partial r} \frac{\partial \phi}{\partial r} + \frac{4}{r} \frac{\partial \psi}{\partial r} \right) \frac{\partial \psi}{\partial r} \right] + \left( -c_{123} \frac{\partial \Lambda}{\partial r} + \frac{c_{123}}{c_{14}} + c_1 + c_3 + 1 \right) \frac{1}{r} \frac{\partial \Lambda}{\partial r} \frac{\partial \psi}{\partial r} \right] \right] \psi = 0. \] (86)

Moreover, using (78) to eliminate the \( \partial \phi/\partial r \) term from (79), we find that
\[ \omega^t = 2c_{123} e^{-2\Phi} (\lambda_2 \psi_1 - \lambda_1 \psi_2) \] (87)

and the reduced symplectic form can be written as
\[ \Omega(\psi_1, \psi_2) = -4\pi c_{123} c_{14} \int_0^\infty dr^r e^{3\lambda - 3\Phi} \left( \frac{\partial \psi_1}{\partial t} \psi_2 - \frac{\partial \psi_2}{\partial t} \psi_1 \right). \] (88)

The form \( W_{\alpha\beta} \) thus defined may be positive or negative definite, depending on the signs of \( c_{123} \) and \( c_{14} \) (recall that we are assuming that \( c_{123} \) and \( c_{14} \) are non-vanishing); however, it is never indefinite. Thus, the symplectic form is in fact of the required form (19), and the...
equations of motion can be put in the form (1). We can therefore write down our variational principle of the form (3); the denominator is

$$
(\psi, \psi) = -4\pi \frac{c_{123} c_{14}}{2 + c_{14}} \int_{0}^{\infty} dr r^4 e^{\Lambda - 3\Phi} \psi^2
$$

and the numerator is

$$
(\psi, T\psi) = -4\pi \frac{c_{123}}{a} \int_{0}^{\infty} dr r^4 e^{-\Phi} \left[ c_{123} \left( \frac{\partial \psi}{\partial r} \right)^2 + \left( c_{123} \frac{\partial \Lambda}{\partial r} - \left( c_{123} - c_2 + 1 \right) \frac{1}{r} \frac{\partial \Phi}{\partial r} - \left( c_{123} + 1 + c_3 + 1 \right) \frac{1}{r} \frac{\partial \Phi}{\partial r} \right) \psi^2 \right].
$$

C. Discussion

We can now examine the properties of this variational principle to determine the stability of Einstein-aether theory. The simplest case to examine is that of flat space. Let us denote the coefficient of $\psi^2$ in (33) by $Z(r)$, i.e.,

$$
Z(r) = c_{123} \frac{\partial \Lambda}{\partial r} \frac{\partial \Phi}{\partial r} - \left( \frac{c_{123} c_{14}}{c_{14}} - c_2 + 1 \right) \frac{1}{r} \frac{\partial \Phi}{\partial r} - \left( \frac{c_{123} + 1 + c_3 + 1}{c_{14}} \right) \frac{1}{r} \frac{\partial \Phi}{\partial r}. \tag{91}
$$

In the case of flat spacetime, $Z(r)$ vanishes, and we are left with

$$
\omega^2 \leq \frac{c_{123}(2 + c_{14})}{c_{14} a} \int_{0}^{\infty} dr r^4 \left( \frac{\partial \psi}{\partial r} \right)^2. \tag{92}
$$

Since both integrands are strictly positive, we conclude that flat space is stable to spherically symmetric perturbations in Einstein-aether theory if and only if

$$
\frac{c_{123}(2 + c_{14})}{c_{14}(c_1 + c_3 + 1)(c_{123} + 2c_2 - 2)} > 0. \tag{93}
$$

This is the same condition found in [11] for the stability of the “trace” wave mode (modulo the aforementioned sign differences). The other stability conditions found in [11] — those corresponding to the “transverse traceless metric” and “transverse aether” wave modes — are incompatible with spherical symmetry, and thus our analysis cannot reproduce these conditions.

More broadly, we can also analyze the stability of the general spherically symmetric solutions described in [24]. These (exact) solutions are described in terms of a parameter $Y$:

$$
r(Y) = r_{min} \frac{Y - Y_-}{Y} \left( \frac{Y - Y_-}{Y - Y_+} \right)^{\frac{1}{\sqrt{Y}}}. \tag{94}
$$

We can then obtain $Z(Y)$ by writing out (33) in terms of these functions of $Y$ (noting that, for example, $\partial \Lambda / \partial r = (\partial \Lambda / \partial Y) / (\partial r / \partial Y)$), and then plot $Z$ and $r$ parametrically. The resulting function is shown in Figure 1 for $c_{123} = \pm c_{14}$ and $c_{14} = -0.1$.

The first thing we see is that in the asymptotic region, the sign of $Z(r)$ is determined by the sign of $c_{123}$. To quantify this, we can obtain an asymptotic expansion of $Z(r)$ as $r \to \infty$, noting that $r \to \infty$ as $Y \to 0$. Doing this, we find that to leading order in $M/r$,

$$
Z(r) = c_{123} \frac{M}{r^3} + \ldots \tag{99}
$$

Thus, the coefficient of the $\psi^2$ term in (33) is, to leading order in the asymptotic region, the same sign as that of the $(\partial \psi / \partial r)^2$ term. We can therefore conclude that for a test function in the “asymptotic region” of one of these solutions (i.e., $r \gg M$), the conditions on the $c_i$’s are the same as those in flat spacetime.\footnote{Note that if the leading-order coefficient in (33) had been different from that of the $(\partial \psi / \partial r)^2$ term in (33), we would have obtained a new constraint on the $c_i$ coefficients; the matching of these two coefficients seems to be coincidental.}

For a realistic
spacetime, of course, there will be some ball of matter in the central region, and the region where the vacuum solution holds will be precisely the region where \( r \gg M \); thus, we can conclude that for a normal star in Einstein-æther theory, the exterior is stable if and only if (93) holds. (Henceforth, we will assume that the \( c_i \) coefficients have been chosen with this constraint in mind, unless otherwise specified.)

We also note that as \( r \) approaches \( r_{\text{min}} \), \( Z(r) \) diverges negatively. One might ask whether a test function in this region (in a spherically symmetric spacetime surrounding a compact object, say) could lead to an instability. We investigated this question numerically using our variational principle; however, our results for \( c_{123} = \pm 0.1 \), respectively; in both cases, \( c_2 = 0 \) and \( c_{14} = -0.1 \). Note that the choice of parameters leading to the dashed plot would lead to instability in the asymptotic region.

![Graph](image)

**FIG. 1:** Representative plots of \( Z(r) \) versus \( r \) for the solutions described in [24]. The solid and dashed lines correspond to \( c_{123} = \pm 0.1 \), respectively; in both cases, \( c_2 = 0 \) and \( c_{14} = -0.1 \). Note that the choice of parameters leading to the dashed plot would lead to instability in the asymptotic region.

---

We conclude that in the \( c_{123} = 0 \) case, the spherically symmetric static solutions of Einstein-æther theory are unstable unless

\[
(c_2 - 1)(c_{14} + 2) > 0, 
\]

D. **The case of \( c_{123} = 0 \)**

The above analysis assumed that \( c_{123} \) was non-vanishing; however, the action originally considered by Jacobson and Mattingly \[6\] used a “Maxwellian” action, i.e., the kinetic terms for the æther field were of the form \( F_{ab}F^{ab} \), corresponding to \( c_3 = -c_1 \) and \( c_2 = 0 \). This case is therefore of some interest.

In the preceding analysis, the perturbational equations of motion (109) and (110) were obtained without any assumptions concerning \( c_{123} \), as was the preconstraint equation (78). We can therefore simply set \( c_{123} \) to zero in these equations and use (78) to solve for \( \partial \phi / \partial r \), as we did in the general case. After application of the background equations of motion, there result the equations

\[
- \frac{2}{r^2} \left( \frac{2}{c_{14}} + 1 \right) \lambda + e^{2\lambda - \Phi} \frac{2}{r^2} (c_2 - 1) \frac{\partial v}{\partial t} = 0 
\]

and

\[
- e^{\Phi} \frac{2}{r} (1 - c_2) \frac{\partial \lambda}{\partial t} - Z_0v = 0, 
\]

where we have defined \( Z_0 \) in terms of the background fields:

\[
Z_0 = \frac{c_{14}}{2} \left( \frac{\partial \Phi}{\partial r} \right)^2 - \frac{2}{c_2} \left( \frac{\partial \Lambda}{\partial r} + \frac{\partial \Phi}{\partial r} \right) + \frac{2 \partial \Lambda}{r \partial r} + \frac{1}{r^2} (e^{2\lambda} - 1). 
\]

We can then combine these two equations to make a single second-order equation for \( v \):

\[
e^{2\lambda - 2\Phi} \frac{c_{14}}{2} (c_2 - 1)^2 \frac{\partial^2 v}{\partial t^2} - Z_0v = 0. 
\]

Since there are no radial derivatives to contend with, the solutions to this equation are all of the form

\[
v(r, t) = f(r) \exp \left( \pm \sqrt{Z_0(r)} \left( \frac{2}{c_{14}} + 1 \right) e^{\Phi - \Lambda} |c_2 - 1| t \right) 
\]

where \( f(r) \) is an arbitrary function of \( r \).

Stability of these solutions will thus depend on the sign of the quantity \( \left( \frac{2}{c_{14}} + 1 \right) Z_0(r) \). We can plot \( Z_0(r) \) parametrically, as was done in the general case; the resulting function, shown in Figure 2, falls off rapidly in the asymptotic region. We can also perform an asymptotic expansion similar to that done in the general case to find the large-\( r \) behaviour of \( Z_0(r) \); the result is that

\[
Z_0(r) \approx (1 - c_2)c_{14} M^2 \frac{r^2}{M^2} + \ldots 
\]

We conclude that in the \( c_{123} = 0 \) case, the spherically symmetric static solutions of Einstein-æther theory are unstable unless

\[
(c_2 - 1)(c_{14} + 2) > 0, 
\]

---

9 We have been assuming a “static æther” throughout this work; however, there also exist solutions in which the æther vector \( \mathbf{u} \) is not globally aligned with the time-translation vector field \( \mathbf{u} \). Thus, even supposing that such “extremely compact stars” are phenomenologically realistic, and that the exterior solutions for such objects are unstable, these “non-static æther” solutions might still be physically viable.
Applying this to our solution (103), we find that for suf-
ficiently large $t$, we have
\[
\frac{\partial^2 \phi}{\partial r^2} \approx \sqrt{(c_2 - 1)^3(c_{14} + 2c_2)} \frac{M_\odot M_\oplus}{R^3 r_\oplus^2} t
\]  
(110)
where $R$ is the radius of Earth’s orbit and $M_\odot$ is the mass of
the Sun. This enhancement to $\partial^2 \phi/\partial r^2$ will lead to an
observable change in the tidal effects due to the Sun’s
gravity; demanding that these remain small relative to
the normal Newtonian tidal effects, i.e., $\partial^2 \phi/\partial r^2 \ll
\partial^2 \Phi/\partial r^2 \approx 2M_\odot/R^3$, we then have
\[
\sqrt{(c_2 - 1)^3(c_{14} + 2c_2)} \approx \frac{r_\oplus^2}{M_\odot t} \approx 1.9 \times 10^{-10}
\]  
(111)
for $t \approx 5 \times 10^9$ years (the approximate age of the Earth.)
We see, therefore, that the requirement that perturba-
tional tidal effects remain small severely constrains
our choices of $c_{14}$ and $c_2$.

V. TeVeS

A. Theory

TeVeS (short for “Tensor-Vector-Scalar”) is a modi-
fied gravity theory proposed by Bekenstein \cite{8} in an at-
tempt to create a fully covariant theory of Milgrom’s
Modified Newtonian Dynamics (MOND). The fields
present in this theory consist of the metric; a vector field
$u^a$ which (as in Einstein-æther theory) is constrained by
a Lagrange multiplier $Q$ to be unit and timelike; and two
scalar fields, $\alpha$ and $\sigma$. The Lagrangian four-form is
\[
\mathcal{L} = (\mathcal{L}_g + \mathcal{L}_v + \mathcal{L}_s + \mathcal{L}_m) \epsilon
\]  
(122)
where $\mathcal{L}_g$ is the usual Einstein-Hilbert action,
\[
\mathcal{L}_g = \frac{1}{16\pi R};
\]  
(133)
$\mathcal{L}_s$ is the “scalar part” of the action,
\[
\mathcal{L}_s = -\frac{1}{2} \sigma^2 (g^{ab} - u^a u^b) \nabla_a \alpha \nabla_b \alpha - \frac{1}{4} \ell^2 \sigma^4 F(\kappa \sigma^2),
\]  
(144)
with $k$ and $\ell$ positive constants of the theory (with
“length dimensions” zero and one, respectively), and
$F(x)$ a free function; $\mathcal{L}_v$ is the vector part of the action,
\[
\mathcal{L}_v = -\frac{K}{32\pi} F_{ab} F^{ab} + Q(u^a u_a + 1),
\]  
(155)
with $K$ a positive dimensionless constant and $F_{ab} = \nabla_a u_b - \nabla_b u_a$; and $\mathcal{L}_m$ the matter action, non-minimally
coupled to the metric:
\[
\mathcal{L}_m \epsilon = \mathcal{L}_{\text{mat}} [A, e^{2\alpha} g^{ab} + 2u^a u^b \sinh(2\alpha)]
\]  
(166)
where $\mathcal{L}_{\text{mat}} [A, g^{ab}]$ would be the minimally coupled mat-
ter Lagrangian for the matter fields $A$. Note that if we

i.e., unless either $c_2 > 1$ and $c_{14} > -2$, or $c_2 < 1$
and $c_{14} < -2$. Note that this excludes the limit of
“small $c_i$” often considered in works such as \cite{7}. In the
“Maxwellian” case, this instability is likely related to the
non-boundedness of the Hamiltonian \cite{27}; however, our
analysis also suggests that non-standard kinetic terms can
in fact stabilize the theory (at least in the time-
like ether case), contrary to the arguments made in that
work.

Finally, we note that even if \cite{106} holds, the perturba-
tional solutions of the form \cite{114} may still be problem-
atic, as the radial gradients of $v$ will grow linearly with
time. To quantify what magnitude of gradients would
be acceptable, we note that in the $c_{123} = 0$ case, there is
only one dynamical degree of freedom; thus, $v$ uniquely
determines $\lambda$ and $\partial \phi/\partial r$. Setting $c_{123} = 0$
in \cite{178} and \cite{79} allows us to solve these equations for $\lambda$ and $\partial \phi/\partial r$; in particular,
\[
\frac{\partial \phi}{\partial r} = - \left[ \frac{2}{r^2} c_{14} e^{2\lambda} + \left( \frac{2}{r} - c_{14} \frac{\partial \Phi}{\partial r} \right) \left( \frac{2}{r} c_2 - c_{14} \frac{\partial \Phi}{\partial r} \right) \right] 
\times \left[ \frac{2}{r^2} c_{14} e^{2\lambda} + \left( \frac{2}{r} - c_{14} \frac{\partial \Phi}{\partial r} \right)^{2^{-1}} e^{2\lambda - \frac{\partial \Phi}{\partial t}} \frac{\partial \Phi}{\partial t},
\]  
(107)
which simplifies in a nearly-flat spacetime to be
\[
\frac{\partial \phi}{\partial r} \approx - \frac{c_{14} + 2c_2}{c_{14} + 2} \frac{\partial \Phi}{\partial t}.
\]  
(108)
Applying this to our solution \cite{108}, we find that for suf-
ficiently large $t$
\[
\frac{\partial^2 \phi}{\partial r^2} \approx \sqrt{(c_2 - 1)^3(c_{14} + 2c_2)} \frac{M_\odot M_\oplus}{R^3 r_\oplus^2} t
\]  
(109)
We can estimate typical scales for $\partial \phi/\partial r$ by looking at
other sources, such as planets. The scale of perturbations
due to the Earth’s gravitational field (at its surface) is
given by $\partial \phi/\partial r \approx M_\odot/r_\odot^2$, where $M_\odot$ and $r_\odot$ are
the mass and radius of the Earth, respectively. Thus, the

\begin{figure}[h]
\centering
\includegraphics[width=\textwidth]{example}
\caption{Plot of $\frac{1}{(\frac{2}{r^2} + 1)} Z_0(r)$, relevant for Einstein-ether theory with $c_{123} = 0$. In this plot, we have chosen $c_{14} = -0.1$ and $c_2 = 0.1$.}
\end{figure}
ignore the scalar and matter portions of the Lagrangian, this Lagrangian is the same as that for Einstein-aether theory \([\mathcal{L}_m = 0]\), with \(c_1 = -c_3 = -K/16\pi\) and \(c_2 = c_4 = 0\). In the present work, we will work exclusively with the “vacuum” \((\mathcal{L}_m = 0)\) theory.

Taking the variation of \((\mathcal{L}_a)\) to obtain the equations of motion and the symplectic current, we find that

\[
\delta\mathcal{L} = ((\mathcal{E}_G)_{ab}\delta g^{ab} + (\mathcal{E}_u)_a\delta u^a + \mathcal{E}_\sigma\delta\alpha + \mathcal{E}_Q\delta Q + \nabla\theta^a)\epsilon, \quad (117)
\]

where

\[
(\mathcal{E}_G)_{ab} = \frac{1}{16\pi} G_{ab} - \frac{1}{2}\sigma^2\nabla_a\nabla_b\alpha - Q u_a u_b
\]

\[
+ \frac{K}{16\pi} \left( 2u_a \nabla^c F_{c}^b - F_{ac} F_{b}^c + \frac{1}{2} g_{ab} F_{cd} F^{cd} \right)
\]

\[
+ \frac{1}{2} g_{ab} \left( \frac{1}{2} \sigma^2 (\nabla^c \alpha \nabla_c \alpha - \dot{\alpha}^2) + \frac{1}{2} \ell^{-2} \sigma^4 F(k\sigma^2) \right), \quad (118a)
\]

\[
(\mathcal{E}_u)_a = \sigma^2 \dot{\alpha} \nabla_a \alpha + \frac{K}{8\pi} \nabla^b F_{ba} + 2 Qu_a, \quad (118b)
\]

\[
\mathcal{E}_\sigma = \nabla_a \left( \sigma^2 (\nabla^a \alpha - u^a \dot{\alpha}) \right), \quad (118c)
\]

\[
\mathcal{E}_Q = -\sigma \left[ \nabla^a \nabla_a \alpha - \dot{\alpha}^2 \right.
\]

\[
+ \ell^{-2} \sigma^2 \left( F(k\sigma^2) + \frac{1}{2} \ell^2 \sigma^4 F'(k\sigma^2) \right), \quad (118d)
\]

and

\[
\theta^a = \theta^a_{\text{Ein}} - \sigma^2 (\nabla^a \alpha - u^a \dot{\alpha})\delta \alpha
\]

\[
+ \frac{K}{8\pi} \left( F_{b}^a \delta u^b + F_{b}^a u_c \delta g^{bc} \right). \quad (119)
\]

In the above, we have defined \(\dot{\alpha} \equiv u^a \nabla_a \alpha\). The remaining equation \(\delta\mathcal{L} = 0\) is identical to that in Einstein-aether theory, \([\mathcal{L}_a]\). For a static solution, in our usual gauge, the background equations of motion become

\[
(\mathcal{E}_G)_{tt} = e^{2\Phi - 2\Lambda} \left[ \frac{1}{16\pi} \left( \frac{2}{r} \frac{\partial \Lambda}{\partial r} + \frac{1}{r^2} (e^{2\Lambda} - 1) \right) \right.
\]

\[
+ \frac{K}{16\pi} \left( -\frac{\partial^2 \Phi}{\partial r^2} + \frac{\partial \Phi}{\partial r} \left( \frac{2}{r} \frac{\partial \Lambda}{\partial r} - \frac{1}{2} \frac{\partial \Phi}{\partial r} \right)^2 \right)
\]

\[
- \frac{1}{4} \sigma^2 \left( \frac{\partial \alpha}{\partial r} \right)^2 - e^{2\Lambda} \frac{1}{8\ell^2} \sigma^4 F(k\sigma^2) \right), \quad (120a)
\]

\[
(\mathcal{E}_G)_{rr} = \frac{1}{16\pi} \left( \frac{2}{r} \frac{\partial \Phi}{\partial r} - \frac{1}{r^2} (e^{2\Lambda} - 1) \right) + \frac{K}{16\pi} \left( \frac{\partial \Phi}{\partial r} \right)^2
\]

\[
- \frac{1}{4} \sigma^2 \left( \frac{\partial \alpha}{\partial r} \right)^2 + e^{2\Lambda} \frac{1}{8\ell^2} \sigma^4 F(k\sigma^2), \quad (120b)
\]

\[
\mathcal{E}_\alpha = \frac{e^{-\Phi - \Lambda}}{r^2} \frac{\partial \theta}{\partial r} \left( \sigma^2 r^2 e^{\Phi - \Lambda} \frac{\partial \alpha}{\partial r} \right), \quad (120c)
\]

and

\[
\mathcal{E}_\sigma = -e^{-2\Lambda} \left( \frac{\partial \alpha}{\partial r} \right)^2 + \frac{1}{k\ell^2} g(k\sigma^2), \quad (120d)
\]

where we have defined \(g(x) = -x F(x) - \frac{1}{2} x^2 F'(x)\) (as in \([8]\)) and used the equation for \(Q\),

\[
Q = \frac{K}{16\pi} e^{-2\Lambda} \left( \frac{\partial^2 \Phi}{\partial r^2} + \frac{\partial \Phi}{\partial r} \left( \frac{\partial \Lambda}{\partial r} - \frac{2}{r} \right) \right), \quad (121)
\]

to simplify. (As in Einstein-aether theory, the background equation \((\mathcal{E}_u)_a = 0\) is satisfied trivially by a static ether.)

The symplectic form for the theory can now be calculated from \([119]\). The result will be essentially the same as that in Einstein-aether theory (with the appropriate values of the \(c_i\)’s), with added terms stemming from the variations of \(\mathcal{L}_a\):

\[
\omega^a = \omega^a_{\text{Ein}} + \omega^a_{\text{vec}} + \omega^a_s \quad (122)
\]

where \(\omega^a_{\text{Ein}}\) is given by \([30]\), \(\omega^a_{\text{vec}}\) is given by \([31]\) with \(c_1 = -c_3 = -K/16\pi\) and \(c_2 = c_4 = 0\), and

\[
\omega^a_s = -\sigma^2 \left[ \nabla^a \alpha - u^a \dot{\alpha} \right] \left( \frac{2}{r} \frac{\partial \sigma}{\partial r} - \frac{1}{2} g_{bc} \delta g^{bc} \right)
\]

\[
+ \delta_1 g^{ab} \nabla_b \alpha + (g^{ab} - u^a u^b) \nabla_b \delta \alpha
\]

\[
- 2u^a \delta_1 u^b \nabla_b \alpha \right) \delta_2 \alpha. \quad (123)
\]

### B. Applying the variational formalism

The next step is to write out the \(t\)-component of \(\omega^a\) in terms of the perturbational variables. We will take our metric perturbations to have the usual form, and our vector perturbation to be of the same form as was used for Einstein-aether theory \([71]\). For the two scalar fields, we define \(\delta \alpha \equiv \beta\) and \(\delta \sigma \equiv \tau\). Calculating \(\omega^a_t\) in terms of these perturbational variables, and using the results of \([75]\), we find that

\[
\omega^t = e^{-2\Phi} \left\{ 2 \sigma^2 \frac{\partial \beta_1}{\partial t} \beta_2 + e^{\Phi} \left[ \frac{K}{8\pi} \left( \frac{\partial \beta_2}{\partial t} - \frac{\partial \beta_1}{\partial r} e^{2\Lambda - \Phi} \frac{\partial \beta_2}{\partial t} + \sigma^2 \frac{\partial \alpha}{\partial r} \right) \right] \right\}
\]

\[
- [1 \leftrightarrow 2]. \quad (124)
\]

We now turn to the question of the constraints. The tensor \(C_{ab}\) is again given by \(C_{ab} = (\mathcal{E}_G)_{ab} + u_a (\mathcal{E}_u)_b\), and thus to first order we have \(\delta C_{tt} = (\mathcal{E}_G)_{tt} + \delta \alpha_\tau (\mathcal{E}_u)_t\) (since \(u^t = 0\) in the background.) Calculating \(\delta C_{tt}\) in terms of our perturbational variables, we find that the
preconstraint equation is

\[ F = r^2 e^{-\Lambda} \left[ \sigma^2 \frac{\partial \alpha}{\partial r} \beta - \frac{1}{8\pi} \left( \frac{2}{r^2} + K \frac{\partial \Phi}{\partial r} \right) \lambda \right. 
\[ \left. + \frac{K}{8\pi} 2\Lambda \Phi \frac{\partial \nu}{\partial r} + \frac{K}{8\pi} \frac{\partial \phi}{\partial r} \right] = 0. \] (125)

In principle, we could now use this equation, together with the equation

\[ 0 = \delta C_{\tau \tau} = e^{2\Lambda} \left( \frac{\sigma^2}{2\ell^2} F(k^2) - \frac{1}{4\pi r^2} \right) \lambda + \sigma \left( \frac{1}{k^2} e^{2\Lambda} y(k - \alpha^2) - \left( \frac{\partial \alpha}{\partial r} \right)^2 \right) \tau + e^{2\Lambda - \frac{K}{8\pi} \frac{\partial \Phi}{\partial r}} \frac{\partial \nu}{\partial r} \frac{\partial \phi}{\partial r} - \sigma \frac{\partial \alpha}{\partial r} \frac{\partial \beta}{\partial r} + \frac{1}{8\pi} \left( \frac{2}{r} + K \frac{\partial \Phi}{\partial r} \right) \frac{\partial \phi}{\partial r} \] (126)

to obtain equations for \( \lambda \) and \( \partial \phi/\partial r \) in terms of the “matter” variables \( v, \beta, \) and \( \tau \). However, it is simpler to pursue a similar tactic to the one we used in the reduction of Einstein-aether theory. To wit, the \( r \)-component of \((\delta \Sigma)_{\tau \tau}\) is

\[ (\delta \Sigma)_{\tau \tau} = \left( 2 e^{2\Lambda} Q + \sigma^2 \left( \frac{\partial \alpha}{\partial r} \right)^2 \right) v + e^{-\phi} \frac{\sigma^2}{2\ell^2} F(k^2) \frac{\partial \alpha}{\partial r} \frac{\partial \beta}{\partial r} \frac{\partial \nu}{\partial r} \frac{\partial \phi}{\partial r} \] (127)

We can use (125) to simplify this equation; the result is

\[ e^\phi \left( 2 e^{2\Lambda} Q + \sigma^2 \left( \frac{\partial \alpha}{\partial r} \right)^2 \right) v = \frac{\partial}{\partial r} \left( \frac{1}{4\pi r} \lambda - 2\sigma^2 \frac{\partial \alpha}{\partial r} \beta \right). \] (128)

We then define the new variable \( \chi \) as

\[ \chi = \frac{1}{4\pi r} \lambda - 2\sigma^2 \frac{\partial \alpha}{\partial r} \beta. \] (129)

The evolution equation for \( \beta \), meanwhile, is given by the equation \( \delta \Sigma_{\phi} = 0 \); in terms of the perturbational fields, it is

\[ -2 e^{2\Lambda - 2\Phi} \frac{\partial \beta}{\partial r}^2 + \frac{\partial \beta}{\partial r}^2 + \left( \frac{2}{r} \frac{\partial \alpha}{\partial r} + \frac{\partial \Phi}{\partial r} - \frac{\partial \Lambda}{\partial r} + \frac{2}{r} \right) \frac{\partial \beta}{\partial r} \]
\[ + \frac{\partial \alpha}{\partial r} \left( \frac{\partial \phi}{\partial r} - \frac{\partial \Lambda}{\partial r} \right) + 2 \frac{\partial \alpha}{\partial r} \frac{\partial \tau}{\sigma \partial r \partial r} \frac{\partial \beta}{\partial r} - \frac{1}{\sigma^2 \partial r \partial r} \frac{\partial \alpha}{\partial r} \frac{\partial \nu}{\partial r} \frac{\partial \phi}{\partial r} \frac{\partial \tau}{\partial r} \] (130)

Finally, the field \( \tau \), being the perturbation of the auxiliary field \( \sigma \), can be solved for algebraically in the equation \( \delta \tau_{\alpha \beta} = 0 \):

\[ \ell^2 \sigma^2 \frac{\partial \phi}{\partial r} \frac{\partial \alpha}{\partial r} \frac{\partial \beta}{\partial r} = 0. \] (131)

We can thus follow the following procedure to write the evolution equations solely in terms of \( \beta \) and \( \chi \): first, we use the preconstraint equation (125) to eliminate the combination \( \partial \beta/\partial r + e^{2\Lambda - \Phi} \partial \nu/\partial t \) in favour of \( \lambda \) and \( \beta \); next, we use (131) to eliminate \( \tau \) and its spatial derivatives; and finally, we use the definition of \( \chi \) (129) to eliminate \( \lambda \) and its derivatives in favour of \( \beta \) and \( \chi \). Applying our procedure to (126) results in an equation depending on \( \partial v/\partial t \) as well as \( \beta, \partial \beta/\partial r, \) and \( \chi \); combining this equation with (128) then yields a second-order evolution equation for \( \chi \) of the form

\[ \frac{\partial^2 \chi}{\partial t^2} = W \chi + V_1 \frac{\partial \beta}{\partial r} + V_4 \beta. \] (132)

where the coefficients \( V_i \) and \( W \) are dependent on the background fields. Similarly, applying this procedure to (130), and using (126) to eliminate a resulting \( \partial v/\partial t \) term, we obtain a second-order equation for \( \beta \):

\[ \frac{\partial^2 \beta}{\partial t^2} = U_1 \frac{\partial \beta}{\partial r} + U_2 \frac{\partial^2 \nu}{\partial r^2} + U_3 \beta + V_1 \frac{\partial \chi}{\partial r} + V_2 \chi, \] (133)

where, again, the \( U_i \)’s and \( V_i \)’s depend on the background fields. Note that (133) has no dependence on the spatial derivatives of \( \chi \), only upon \( \chi \) itself.

We also need to reduce the symplectic form and express it in terms of \( \beta \) and \( \chi \). Applying the preconstraint equation (125) to (124), and using the definition of \( \chi \) along with (128), we find that

\[ \Omega = 4\pi \int dr r^2 e^{-\Lambda} \left[ \frac{1}{H} \frac{\partial \chi}{\partial t} \right] \left[ 2\sigma^2 \frac{\partial \beta}{\partial r} \right] \right] \] (134)

where

\[ H = - \left( 2 e^{2\Lambda} Q + \sigma^2 \left( \frac{\partial \alpha}{\partial r} \right)^2 \right). \] (135)

We can also apply the background equations of motion, along with the Equation (121), to this quantity to obtain

\[ H = \frac{1}{4\pi r^2} \left( \frac{\partial H}{\partial r} + \frac{\partial \Phi}{\partial r} \right) - 2\sigma^2 \left( \frac{\partial \alpha}{\partial r} \right)^2. \] (136)

For a variational principle to exist for this theory, the form \( W_{\alpha \beta} \) defined by (133) must be positive definite. In our case, this means that the coefficients of both \((\partial \chi/\partial t) \chi_2 \) and \((\partial \beta_1/\partial t) \beta_2 \) in the integrand of (134) must be positive for the background solution about which we are perturbing. While the coefficient for the latter term is obviously always positive, the situation for the former coefficient (namely, \( H \)) is not so clear. To address this issue, we need to know the properties of the spherically symmetric static background solutions of TeVeS. These solutions (with a “static aether”) are described in [23]. In our gauge, they are most simply described in
This is defined by an integral over the central mass distribution $\Phi(r)$, while $
abla$ non-negative. Finally, the constant $z$ approximated. Calculating this function is strictly negative. In fact, in the $z \to \infty$ limit, we have $r/z = c + O(z_c/z)$, and so we can take $r \approx z$ to a good approximation. Calculating $H(r)$ in terms of $\Phi(r)$ and $\Lambda(r)$, we find that as $r \to \infty$,

$$H(r) \approx -\frac{1}{16\pi} \left( \frac{k}{m^2} + \frac{K}{2} \right) \frac{1}{r^4}$$

which is negative for any positive choice of $k$ and $K$. Therefore, it is not possible to straightforwardly apply the variational principle to TeVeS, since the quadratic form defined by $W_{\alpha\beta}$ is indefinite and thus cannot used as an inner product.\(^{11}\)

C. WKB analysis

While we cannot derive a variational principle with TeVeS, the fact that we have been able to \textquotedblleft reduce\textquotedblright the equations of motion to an unconstrained form still allows us to analyse the stability of its spherically symmetric solutions. Let us consider a WKB ansatz, of the form

$$\begin{bmatrix} \beta(r) \\ \chi(r) \end{bmatrix} = e^{i(\omega(r)t + \kappa r)} \begin{bmatrix} f_\beta(r) \\ f_\chi(r) \end{bmatrix}$$

with $\kappa$ very large compared to the scale of variation of the background functions $U$, $V$, and $W$. We will further choose the the functions $f_\beta(r)$ and $\omega(r)$ are chosen to be \textquotedblleft slowly varying\textquotedblright relative to the scale defined by $\kappa$, i.e.

$$\frac{1}{f_\beta(r)} \frac{\partial f_\beta(r)}{\partial r} \ll \kappa, \quad \frac{1}{\omega(r)} \frac{\partial \omega(r)}{\partial r} \ll \kappa.$$

Under this assumption, we will then have

$$\frac{\partial}{\partial r} \left( e^{i(\omega(r)t + \kappa r)} f_\beta(r) \right) \approx i\kappa e^{i(\omega(r)t + \kappa r)} f_\beta(r).$$

Now let us apply the time-evolution operator $\mathcal{T}$ implicitly defined by (133) and (132) to our ansatz (143). We see that for sufficiently large $\kappa$, the highest-derivative terms will dominate the lower-derivative terms. Thus, to a good approximation we have

$$\mathcal{T} \begin{bmatrix} \beta(r) \\ \chi(r) \end{bmatrix} \approx \begin{bmatrix} -\kappa^2 U_1(r) \\ \kappa V_3(r) \end{bmatrix} \begin{bmatrix} f_\beta(r) \\ f_\chi(r) \end{bmatrix} e^{i(\omega(r)t + \kappa r)}.$$

Then, in the limit of large $\kappa$, our ansatz (143) will be an approximate eigenvector of $\mathcal{T}$ if there exist a $f_\beta(r)$ and $f_\chi(r)$ such that

$$-\omega^2(r) \begin{bmatrix} f_\beta(r) \\ f_\chi(r) \end{bmatrix} = \begin{bmatrix} -\kappa^2 U_1(r) \\ \kappa V_3(r) \end{bmatrix} \begin{bmatrix} f_\beta(r) \\ f_\chi(r) \end{bmatrix}.$$

In other words, in the limit of large $\kappa$, the problem of finding modes of $\mathcal{T}$ is a simple two-dimensional eigenvalue problem where the eigenvalues are functions of $r$. In this limit, the eigenvalues of this matrix are (to leading order in $\kappa$)

$$\omega^2(r) \approx \left\{ \kappa^2 U_1(r), -W(r) + \frac{V_1(r)V_3(r)}{U_1(r)} \right\}.$$

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\(^{10}\) This is the \textquotedblleft radial coordinate\textquotedblright in isotropic spherical coordinates, as used in [23].

\(^{11}\) Note that this also implies that the \textquotedblleft perturbational Hamiltonian\textquotedblright of TeVeS, as defined in equation (54) of [1], has an indefinite kinetic term.
It is not difficult to verify that for sufficiently large $\kappa$, our assumptions for the ansatz \(144\) are satisfied.

It remains to write out the functions $U_1(r)$, $V_1(r)$, $V_3(r)$, and $W(r)$ in terms of the background functions. These can be shown to be:

$$U_1 = \frac{1}{2} e^{2\Phi - 2\Lambda} \left( 1 + e^{-2\Lambda} \frac{2\ell^2 \left( \frac{\partial \alpha}{\partial r} \right)^2}{\sigma^2 y'(k\sigma^2)} \right) , \quad (149)$$

$$V_1 = -\frac{r}{4} e^{2\Phi - 2\Lambda} \frac{\partial \alpha}{\partial r} \left( 1 + e^{-2\Lambda} \frac{2\ell^2 \left( \frac{\partial \alpha}{\partial r} \right)^2}{\sigma^2 y'(k\sigma^2)} \right) , \quad (150)$$

$$V_3 = 16\pi r e^{2\Phi - 2\Lambda} \frac{\partial \alpha}{\partial r} \left( 8\pi \sigma^2 \left( \frac{\partial \alpha}{\partial r} \right)^2 - \frac{1}{r} \left( \frac{\partial \Lambda}{\partial r} + \frac{\partial \Phi}{\partial r} \right) \right) , \quad (151)$$

and

$$W = r^2 e^{2\Phi - 2\Lambda} \left( \frac{2}{K} - 1 - 8\pi \sigma^2 \left( \frac{\partial \alpha}{\partial r} \right)^2 \right) \times \left( 8\pi \sigma^2 \left( \frac{\partial \alpha}{\partial r} \right)^2 - \frac{1}{r} \left( \frac{\partial \Lambda}{\partial r} + \frac{\partial \Phi}{\partial r} \right) \right) . \quad (152)$$

This implies that the eigenvalues of the matrix in \(141\) are

$$\omega^2(r) \approx \frac{K^2}{2} e^{2\Phi - 2\Lambda} \left( 1 + e^{-2\Lambda} \frac{2\ell^2 \left( \frac{\partial \alpha}{\partial r} \right)^2}{\sigma^2 y'(k\sigma^2)} \right) \quad (153)$$

and

$$\omega^2(r) \approx e^{2\Phi - 2\Lambda} \left( \frac{2}{K} - 1 \right) \times \left( \frac{1}{r} \left( \frac{\partial \Lambda}{\partial r} + \frac{\partial \Phi}{\partial r} \right) - 8\pi \sigma^2 \left( \frac{\partial \alpha}{\partial r} \right)^2 \right) . \quad (154)$$

We can see that this first eigenvalue \(153\) is always positive as long as $y'(x) > 0$; indeed, the choice of $y(x)$ made in [8] does satisfy this inequality. The second eigenvalue \(154\), however, is just

$$\omega^2(r) = 4\pi e^{2\Phi - 2\Lambda} \left( \frac{2}{K} - 1 \right) H(r) . \quad (155)$$

Since $H(r)$ is always negative, we conclude that this second mode is unstable for $0 < K < 2$. Further, for a spherically symmetric solution outside a Newtonian star, the approximation \(142\) is valid; thus, to lowest nonvanishing order in $r^{-1}$, we have

$$\omega^2(r) \approx -\frac{1}{4} \left( \frac{2}{K} - 1 \right) \left( \frac{k}{\pi} m_s^2 + \frac{K}{2} z_g^2 \right) \frac{1}{r^4} . \quad (156)$$

### D. Discussion

The above result for $\omega^2(r)$ implies that for stability of Newtonian solutions, we must have $K > 2$, ruling out the range of parameters originally considered in [8]. To estimate the time scale of the instability when $K < 2$, we first note that unless $k$ and $K$ are much larger than the ratio of the star’s radius $R$ to its Schwarzchild radius $m_s$, the parameters $m_s$, $m_s$, and $z_g$ are all approximately equal, differing by terms of $\mathcal{O}(m_s/R)$ [8]. In particular, if $k$ and $K$ are of the same order and $K \neq 2$, we find that the timescale of this instability will be on the order of $10^8$ seconds—approximately two weeks—for points near the surface of the Sun.

Note also that the effective gravitational potential for nonrelativistic motion is

$$\Phi = \Xi \Phi_N + \alpha , \quad (157)$$

where

$$\Xi = e^{-2\alpha_c}(1 - K/2)^{-1} , \quad (158)$$

$\Phi_N$ is the Newtonian gravitational potential, and $\alpha_c$ is the asymptotic value of $\alpha$, determined by the cosmological boundary conditions [29]. Expanding the asymptotic behaviour of $\Phi_N$ and $\alpha$ about infinity, we find that

$$\Phi \approx \alpha_c - e^{-2\alpha_c} \left( \frac{1}{1 - K/2} + \frac{k}{4\pi} \right) \frac{m_s}{r} \quad (159)$$

where in the weak-field limit, we have

$$m_s \approx 4\pi \int r^2 \rho \quad (160)$$

and $m_s \approx e^{-2\alpha_c} m_s$. The coefficient in parentheses in \(159\) can be thought of as renormalizing the gravitational constant $G$ by some overall factor; the first term in parentheses comes from the usual perturbations to the “physical metric” component $\tilde{g}_{tt}$, while the second term comes from variation of the scalar field $\alpha$. However, if $K > 2$, this would make the metric contribution to the “effective gravitational constant” negative. This could be remedied by choosing $k > 8\pi/(K - 2)$. However, it is unclear whether the theory, in this parameter regime, would still be experimentally viable; for example, Big Bang nucleosynthesis constraints require that $k < 0.75$, and the CMB power spectrum also constrains the viable regions of parameter space [30].

\[12\] We also note that if the gravitational effects of the scalar are dominant over those of the tensor, the solutions of TeVeS would greatly resemble those of the so-called “stratified theories” with time-orthogonal space slices [31]. As these are ruled out experimentally via geophysical experiments, we could not make $k$ or $K$ too large without running afool of these experimental constraints.
We also note that the eigenvector corresponding to the unstable mode of $\mathcal{T}$ will satisfy (to leading order)

$$i\kappa f_\beta + \frac{V_1}{U_1} f_\chi \approx 0$$  \hspace{1cm} (161)

or, in the limit of large $\kappa$, $f_\beta \approx 0$. In other words, this unstable mode should manifest itself in growth of the radial component of the vector field $u^\alpha$ rather than growth of the scalar $\alpha$ (cf. (123)). An instability of the vector field was also found in [32] by Dodelson and Liguori. However, it seems unlikely that this is the same instability found in [32] by Dodelson and Liguori. How- ever, the instability found in [32] manifests itself only in a cosmological context, not a Newtonian-gravity context; since the cosmological solutions of TeVeS are in a very real sense separate from the Newtonian solutions (existing on two different branches of the function $y(x)$), it is difficult to draw a direct correspondence between the stability properties of these two types of solutions.

Second, the instability found in [32] manifests itself only in the limit of a matter-dominated Universe; the instability we have found exists in vacuo. Third, Dodelson and Liguori’s instability requires $K$ to be sufficiently small relative to $k$, while our instability is present for all $k > 0$ and $0 < K < 2$

Finally, it is perhaps notable that the vector field $u^\alpha$ in TeVeS has “Maxwellian” kinetic terms, which we found in Section [XX] to be unstable in the context of Einstein-æther theory. It is possible, then, that this instability could be cured via a more general kinetic term for the vector field.

### APPENDIX A: COEFFICIENTS FOR THE $f(R)$ VARIATIONAL PRINCIPLE

The quantity $(\psi, \mathcal{T}\psi)$ for $f(R)$ gravity can be put in the form

$$\langle \psi, \mathcal{T}\psi \rangle = 4\pi \int dr \left[ C_1 \left( \frac{\partial b}{\partial r} \right)^2 + C_2 \left( \frac{\partial c}{\partial r} \right)^2 + C_3 \left( \frac{\partial \zeta}{\partial r} - b \frac{\partial \zeta}{\partial r} \right) + C_4 \nu^2 + C_5 c^2 + 2C_6 \nu c \right]$$  \hspace{1cm} (A1)

where the $C_i$ coefficients are given by

$$C_1 = 6 e^{\Phi - \Lambda} S^{-2} f'(\alpha),$$  \hspace{1cm} (A2)

$$C_2 = 8\pi r^2 e^{\Lambda + \Phi} \nu^2,$$  \hspace{1cm} (A3)

$$C_3 = 16\pi r^2 e^{\Phi + \Lambda} S^{-1} \left( \frac{3}{r^2} S^{-1} f'(\alpha) (\nu - \nu^2) - \nu \left( \frac{\partial \Phi}{\partial r} + \frac{2}{r} \right) \right),$$  \hspace{1cm} (A4)

$$C_4 = r e^{\Phi - \Lambda} \left\{ \frac{6}{r} S^{-1} \left( \frac{2}{r} \frac{\partial \Lambda}{\partial r} - \frac{\partial \Phi}{\partial r} + \frac{2}{r} \right) \right. + S^{-2} \left[ \frac{2}{r^2} \frac{e^{2\Lambda}}{f'(\alpha)} - e^{2\Lambda} \left( \frac{48\pi}{r} \frac{\partial \Phi}{\partial r} + f(\alpha) - 16\pi \zeta \right) \left( \frac{\partial \Lambda}{\partial r} - \frac{\partial \Phi}{\partial r} + \frac{3}{r} \right) \right] + 2f'(\alpha) \left( e^{2\Lambda} - \frac{\partial \Phi}{\partial r} + \frac{5}{r} \right) \right. $$

$$\left. \left. \left. - \frac{3}{r} \left( \frac{1}{r^2} + 7e^{2\Lambda} - 1 \right) + 3 \left( \frac{\partial \Lambda}{\partial r} \right)^2 + \frac{\partial \Lambda}{\partial r} \left( \frac{1}{r} (e^{2\Lambda} - 3) - \frac{\partial \Phi}{\partial r} (e^{2\Lambda} + 7) - \frac{\partial^2 \Phi}{\partial r^2} \right) \right) \right] \right. $$

$$\left. + \frac{6}{r} S^{-3} f'(\alpha) \left[ -8\pi e^{2\Lambda} \frac{\partial \Phi}{\partial r} (\frac{\partial \nu}{\partial r^2}) + 2 \frac{e^{2\Lambda}}{r} f'(\alpha) + 2f'(\alpha) \left( -e^{2\Lambda} + \frac{3}{r} \frac{\partial \Lambda}{\partial r} - \frac{\partial \Phi}{\partial r} + \frac{1}{r} (e^{2\Lambda} - 1) \right) \right] \right. $$

$$\left. \left. \left. + 8\pi e^{2\Lambda} \nu \left( \frac{\partial \Lambda}{\partial r} + \frac{\partial \Phi}{\partial r} + \frac{8}{r} \right) \right) \right] \right. $$

$$\left. - S^{-4} \frac{12}{r} e^{2\Lambda} f'(\alpha) \left( e^{2\Lambda} (8\pi \frac{\partial \nu}{\partial r})^2 + 24\pi \frac{r}{r^2} f'(\alpha) (3\nu - \nu^2) \right) \right\}, \hspace{1cm} (A5)$$

$$C_5 = 8\pi r^2 e^{\Lambda + \Phi} S^{-2} \left\{ \frac{16\pi e^{2\Lambda}}{r} \nu \left( 8\pi e^{2\Lambda} \frac{\partial \nu}{\partial r} + \frac{3}{r^2} f'(\alpha) (\nu^2 + \nu) \right) $$

$$+ 8\pi e^{2\Lambda} S \nu \left( \frac{\partial \nu}{\partial r} - \frac{4}{r} \right) + \nu \left( \left( \frac{\partial \nu}{\partial r} - \frac{4}{r} \right) \frac{\partial \nu}{\partial r} - \frac{4}{r^2} \right) $$

$$- \nu \left( \left( \frac{\partial \nu}{\partial \rho} \right)^2 + 2 \frac{\partial \Lambda}{\partial \rho} + \frac{1}{r} \frac{\partial \Phi}{\partial \rho} + \frac{1}{r^2} (e^{2\Lambda} - 1) - e^{2\Lambda} \frac{\partial \nu}{\partial \rho} + \frac{2}{r} \right) \right\}, \hspace{1cm} (A6)$$
and

\[ C_0 = 8\pi c^{\Lambda+\Phi} r^2 S^{-3} \left\{ S^2 \left[ \left( \frac{\varphi'}{\varphi''} \right)^2 \frac{\partial \Phi}{\partial r} + \left( \frac{\partial \Phi}{\partial r} + \frac{2}{r} \right) \right] + g' \nu \left( \frac{1}{r^2} + \frac{4}{r} \frac{\partial \Phi}{\partial r} + \left( \frac{\partial \Phi}{\partial r} + \frac{2}{r} \right)^2 + \frac{\partial \Lambda}{\partial r} \left( \frac{\partial \Phi}{\partial r} + \frac{2}{r} \right) - \frac{3}{r^2} g'^{\nu} \nu \right] \\
+ \frac{12}{r^2} f'(\alpha) \left( 8\pi g'' \varphi + \frac{3}{r^2} \left( \nu g' + \nu'' \nu \right) \right) \right\} - S \left[ e^{2\Lambda} g' \nu \left( \frac{1}{r} (f(\alpha) - 16\pi g) + 8\pi g' \nu \frac{\partial \Phi}{\partial r} + \frac{2}{r} \right) \right] + \frac{1}{r} f'(\alpha) \left( 3 \frac{\partial \Phi}{\partial r} \left( \varphi' \nu \right)^2 - \frac{3}{r} g'^{\nu} \nu \left( \frac{\partial \Lambda}{\partial r} + \frac{\partial \Phi}{\partial r} + \frac{10}{r} \right) \right) + g' \nu \left( 3 \left( \frac{2}{r} (e^{2\Lambda} - 4) \right) + \frac{\partial \Lambda}{\partial r} + \frac{\partial \Phi}{\partial r} \left( \nu' \nu - 2 \right) - 2e^{2\Lambda} \right) \right], \tag{A7} \]

**APPENDIX B: SYMPLECTIC CURRENT FOR EINSTEIN-Æther THEORY**

The symplectic current for vacuum Einstein-æther theory is given by \( \omega^a = \omega^a_{\text{Ein}} + \omega^a_{\text{vec}} \), where \( \omega^a_{\text{Ein}} \) is the standard Einstein symplectic form (given by \( [30] \)), and

\[
\omega^a_{\text{vec}} = (M_{(1)})^a_{bcde} \delta g^{bc} \nabla_g \delta g^{ef} + (M_{(2)})^a_{bcde} \delta g^{bc} \delta g^{de} + (M_{(3)})^a_{bcde} \delta g^{bc} \nabla_g \delta_1 u^e + (M_{(4)})^a_{bcde} \delta g^{bc} \delta_1 u^d \\
+ (M_{(5)})^a_{bcde} \delta g^{bc} \nabla_g \delta_1 u^c + (M_{(6)})^a_{bcde} \delta g^{bc} \delta_1 u^c + (M_{(7)})^a_{bcde} \delta g^{bc} \delta_1 u^c - [1 \leftrightarrow 2], \tag{B1} \]

where \( M_{(i)} \) tensors are given by

\[
(M_{(1)})^a_{bcde} = c_1 \left( g^{ad} g_{be} u_c u_f - \delta^a_{e} \epsilon^d b u_c u_f + \frac{1}{2} u^a u^d g_{be} g_{cf} \right) + \frac{1}{2} c_2 u^a u^d g_{bc} g_{ef} \\
+ c_3 \left( \delta^a_{e} \delta^d b u_c u_f - g^{ad} g_{be} u_c u_f + \frac{1}{2} u^a u^d g_{be} g_{cf} \right) - c_4 \left( 2 u^a u^d g_{be} - \delta^a_{e} u^d u_b - u^a u_e \delta^d b + \frac{1}{2} g^{ad} u_b u_e \right) u_c u_f, \tag{B2} \]

\[
(M_{(2)})^a_{bcde} = (c_3 - c_1) \left( \nabla_g u^a g_{ce} u_d + \delta^a_{e} \delta^d b u_c u_f + \frac{1}{2} u^a u^d g_{be} g_{ef} \right) - c_4 \left( 2 u^a u^d g_{be} - \delta^a_{e} u^d u_b - u^a u_e \delta^d b + \frac{1}{2} g^{ad} u_b u_e \right) u_c u_f, \tag{B3} \]

\[
(M_{(3)})^a_{bcde} = c_1 \left( \delta^a_{e} \delta^d b u_c u_f - g^{ad} g_{be} u_c - u^a \delta^d g_{ce} \right) - c_2 u^a g_{bc} \delta^d e \\
+ c_3 \left( g^{ad} g_{be} u_c - \delta^a_{e} \delta^d b u_c - u^a \delta^d g_{ce} \right) - c_4 \left( \delta^a_{e} u_c - 2 u^a g_{ce} \right) u_b u_d, \tag{B4} \]

\[
(M_{(4)})^a_{bcde} = c_1 \left( \nabla_g u^a g_{cd} + \nabla^a u_b g_{cd} + \nabla^a u_d g_{be} - \delta^a_{e} \delta^d b u_c - 2 \delta^a_{d} \nabla_g \delta b u_d \right) \\
+ c_3 \left( \nabla^a u_d g_{bc} - \nabla_g u^a g_{cd} - \delta^a_{e} \delta^d b u_c + \nabla d u^a g_{be} \right) \\
- c_4 \left( 2 u^a u_b g_{cd} + u_b u_c \nabla g^a + 2 u^a u_b \nabla_g u^d - 2 u^a u_b \nabla_d u_c + u^a g_{be} u_c \right), \tag{B5} \]

\[
(M_{(5)})^a_{bcde} = K^a_{db} u_e - K^a_{ac} b u_e - K^a_{be} u_c, \tag{B6} \]

\[
(M_{(6)})^a_{bcde} = 2 K^a_{bd}, \tag{B7} \]

and

\[
(M_{(7)})^a_{bc} = -2 c_4 (\delta^a_{e} u_b + u^a \nabla g_{ce} u_b). \tag{B8} \]
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