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Published in:
Astronomy & Astrophysics

DOI:
10.1051/0004-6361/202244661

Publication date:
2022

Document version
Publisher's PDF, also known as Version of record

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Citation for published version (APA):
Sillassen, N. B., Jin, S., Magdis, G. E., Daddi, E., Weaver, J. R., Gobat, R., Kokorev, V., Valentino, F., Finoguenov, A., Shuntov, M., Gomez-Guijarro, C., Coogan, R., Greve, T. R., Toft, S., & Sese, D. B. (2022). A galaxy group candidate at z approximate to 3.7 in the COSMOS field. Astronomy & Astrophysics, 665, [L7]. https://doi.org/10.1051/0004-6361/202244661

Download date: 13. jul., 2024
A&A 665, L7 (2022)
https://doi.org/10.1051/0004-6361/202244661
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LETTER TO THE EDITOR

A galaxy group candidate at $z \approx 3.7$ in the COSMOS field

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Received 2 August 2022 / Accepted 2 September 2022

ABSTRACT

We report a galaxy group candidate HPC1001 at $z \approx 3.7$ in the COSMOS field. This structure was selected as a high galaxy overdensity at $z > 3$ in the COSMOS2020 catalog. It contains ten candidate members, of which eight are assembled in a $10'' \times 10''$ area with the highest sky density among known protoclusters and groups at $z > 3$. Four out of ten sources were also detected at 1.2 mm with Atacama Large Millimeter Array continuum observations. Photometric redshifts, measured by four independent methods, fall within a narrow range of $3.5 < z < 3.9$ and with a weighted average of $z = 3.65 \pm 0.07$. The integrated far-IR-to-radio spectral energy distribution yields a total UV and IR star formation rate SFR $\approx 900 M_{\odot}$ yr$^{-1}$. We also estimated a halo mass of $\sim 10^{15} M_{\odot}$ for the structure, which at this redshift is consistent with potential cold gas inflow. Remarkably, the most massive member has a specific star formation rate and dust to stellar mass ratio of $M_{\text{gas}}/M_{\ast}$ that are both significantly lower than that of star-forming galaxies at this redshift, suggesting that HPC1001 could be a $z \approx 3.7$ galaxy group in maturing phase. If confirmed, this would be the earliest structure in maturing phase to date, and an ideal laboratory to study the formation of the earliest quiescent galaxies as well as cold gas accretion in dense environments.

Key words. Galaxy: evolution – galaxies: high-redshift – submillimeter: galaxies – galaxies: clusters: general

1. Introduction

In the last two decades, massive galaxy protoclusters (see Overzier 2016 for a review) and groups have been discovered out to $z \sim 5$ (e.g., Pentericci et al. 2000; Daddi et al. 2009, 2017; Capak et al. 2011; Walter et al. 2012; Dannerbauer et al. 2014; Cucciati et al. 2018; Oteo et al. 2018; Miller et al. 2018). Theoretical work predicts that they play an important role in the early universe, significantly contributing to the cosmic star formation rate density (20–50%) and forming most of their present-day stellar mass at $z > 1.5$ (Chiang et al. 2017). In a simplified picture, galaxy protoclusters and groups are speculated to experience a “growth phase” at $z > 3$ (Shimakawa et al. 2018), forming their stellar masses via vigorous star formation and growing the total gas reservoirs by accreting cold gas streams in hot media (Dekel et al. 2013; Daddi et al. 2021, 2022). Later on, the core region of the cluster collapses, active galactic nucleus (AGN) activity takes place in cluster galaxies, and star formation declines – the so-called maturing phase (Shimakawa et al. 2018). Eventually, the structure is totally collapsed, and it evolves into a virialized cluster or group of galaxies in the local Universe, is depleted of the cold gas, and has negligible star formation (e.g., Ata et al. 2022).

However, the evolutionary path of these massive structures is still under debate. One of the open questions is when and where the earliest maturing phase is taking place. The virialization of the dark matter halo and the quiescence of cluster members are two critical indicators of the maturing phase. Finding the earliest collapsed dark matter halos and the first generation of quenching cluster members is key to answering this question.

In recent years, compact structures have been found with collapsed dark matter halos and cluster members that form the red sequence of passive galaxies at $z = 2–2.5$ (e.g., Wang et al. 2016; Willis et al. 2020), indicating that their maturing phase took place at $z > 2.5$. SPT2349-56 at $z = 4.3$ is so far the most distant, approximately virialized structure with a dark matter halo mass $M_{\text{halo}} \sim 10^{13} M_{\odot}$ (Miller et al. 2018; Hill et al. 2020) and a promising progenitor of a massive cluster with $M_{\text{halo}} \sim 10^{15} M_{\odot}$ at $z = 0$. However, SPT2349-56 is super-starbursting (star formation rate SFR $\gtrsim 6000 M_{\odot}$ yr$^{-1}$) with rich gas reservoirs ($M_{\text{gas}} \gtrsim 6 \times 10^{11} M_{\odot}$) and it is unlikely that it harbors quenched galaxies. On the other hand, quiescent members in other dense environments have already been identified at $z \sim 3$ (Kubo et al. 2021; Kalita et al. 2021), indicating that some...
structures start to decrease their star formation activity and enter into the maturing phase between \( z = 3 \) and 4.

In this Letter we report a galaxy group candidate HPC1001 at \( z \approx 3.7 \) in maturing phase. We adopt cosmology \( H_0 = 73, \Omega_M = 0.27 \), and \( \Lambda_0 = 0.73 \) as well as a Chabrier initial mass function (Chabrier 2003).

2. Selection, data, and measurements

2.1. Selection

Using GALCLUSTER\(^1\), which is publicly available software designed to automatically search for galaxy overdensities, we mapped the distance of the fifth neighborhood \( \Sigma_{\text{inh}} \) for \( i^- \) and \( K_s^- \) detected sources (\( S/N_i > 5 \) and \( K_s < 25 \) AB mag) at \( z_{\text{phot}} > 3 \) in the COSMOS2020 catalog (Weaver et al. 2022). HPC1001, which is centered at RA 150.4656\(^\circ\) and Dec 2.6359\(^\circ\), is the highest overdensity with \( \Sigma_{\text{inh}} \) that is 6.8\( \sigma \) larger than the mean density of the COSMOS2020 catalog at \( z_{\text{phot}} > 3 \). As shown in Fig. 1, HPC1001 hosts a compact group of galaxies with \( z_{\text{phot}} \approx 3.7 \). Remarkably, the central region contains eight galaxies in a 10\( ^{\prime\prime} \times 10^{\prime\prime} \) area (70\( ^{\prime\prime} \times 70^{\prime\prime} \) kpc\(^2\), Fig. 1-right). Comparing it to protoclusters in literature, HPC1001 has the highest sky density of galaxies at \( z > 3 \) (Fig. A.1). On a large scale, HPC1001 is surrounded by near-IR- (NIR-) detected galaxies (Fig. 1-left) at a similar redshift 3.3 < \( z < 3.9 \) (475 cMpc), and it is hosted by an overdensity of submillimeter galaxies (Fig. A.2). The density of submillimeter galaxies around HPC1001 is \( \sim 0.5 \) arcmin\(^{-2} \) for \( S_{850 \mu m} > 5 \) mJy sources, which is comparable with that in the Spiderweb protocluster at \( z = 2.16 \) (Dannerbauer et al. 2014).

2.2. Optical+NIR SEDs and photometric redshifts

We identified ten sources as candidate members within the virial radius (see Sect. 3.1 for the calculation), of which eight are detected in the COSMOS2020 catalog (Weaver et al. 2022; Fig. 1). COSMOS2020 uses two different methods for obtaining photometry: the aperture photometry (Classic) and The Farmer and two algorithms, LEPHARE and EAZY (Ilbert et al. 2006; Brammer et al. 2008), for the spectral energy distribution (SED) fitting of each of the two versions of photometry. Thus, for the eight sources that are included in the COSMOS2020 catalog, we have four independent photo-\( z \) estimates that fall within a narrow range 3.5 < \( z < 3.9 \) (306 cMpc, Fig. B.1) and for each source are consistent within uncertainties. Six of these eight galaxies are also included in the COSMOS2015 catalog (Laigle et al. 2016), with photo-\( z \) consistent with those derived in our work based on the COSMOS2020 photometry (Fig. B.1). We note though that the COSMOS2020 Classic catalog (and the COSMOS2015 catalog) is using aperture photometry of 2\( ^{\prime\prime} \) (3\( ^{\prime\prime} \)), which is severely affected by blending in crowded regions such as HPC1001. On the other hand, Farmer photometry is extracted by simultaneous modeling of the profiles of all galaxies within the group, making it more robust against blending. Furthermore, we calculated the systematic redshift bias at 3.5 < \( z < 4 \) by cross-matching the COSMOS spec-\( z \) catalog (Salvato et al., in prep.) to the Farmer LEPHARE photo-\( z \), and found it to be negligible, that is to say median \( (z_{\text{phot}} - z_{\text{spec}})/(1 + z_{\text{spec}}) = 1.3 \times 10^{-3} \). Therefore, we chose to adopt the COSMOS2020/Farmer – LEPHARE combination as the benchmark for our analysis, in which the average photo-\( z \) uncertainty of this sample is 0.2.

\(^1\) https://github.com/Nikolaj-B-Sillassen/galcluster

In Fig. 2 we present the optical and NIR SEDs for the eight sources, and list the derived photo-\( z \), stellar mass estimates in Table 1. Apart from a galaxy template, we also performed a fit assuming a quasar template to examine the systematics of the galaxy template solutions (but not to serve as a bona fide AGN classification). A galaxy template provides the best fit for seven out of eight sources, ensuring the fidelity of the derived stellar masses. The SEDs of HPC1001.d and f show a blue excess to the galaxy template solution, suggesting potential AGN activity. As a sanity check, we also compared the redshift probability distribution functions (PDFs) as derived based on Bayesian and \( \chi^2 \) methods and were found to be consistent (Fig. 2).

The remaining two sources, HPC1001.a and h, are not included in COSMOS2020 catalog, as they are only detected in \( K_s \) while their signal-to-noise ratios (S/Ns) are lower than the detection criteria in the \( JHK_s \)-stacked image adopted by the COSMOS2020 analysis. However, HPC1001.h is included in the COSMOS2015 catalog with \( z_{\text{phot}} = 3.75 \pm 0.04 \) (Table 1).

In order to properly deblend HPC1001, we repeated the Farmer photometry this time by adding HPC1001.a and h in the prior catalog (at their respective ALMA position for a and \( K_s \) position for h), and by simultaneously fitting all sources in the \( K_s \) and IRAC images. The new photometry was used to derive the properties of HPC1001.a while it remained virtually unchanged for the eight sources that were included in the original COSMOS2020 catalog. The SED of HPC1001.h is highly uncertain due to its nondetection at wavelengths shorter than \( K_s \) and its faintness in IRAC (\( M_{\text{AB}} > 28 \)). While in Table 1, we list its parameters as derived in the COSMOS2015 catalog, we caution that its photo-\( z \) and \( M_\odot \) are only indicative and their uncertainties could be underestimated.

2.3. FIR, (sub)millimeter, and radio SED

HPC1001 is covered by Atacama Large Millimeter Array (ALMA) Band 6 continuum imaging (Fig. 1-right), as part of the project ID: 2013.1.00034.S (PI: N. Scoville). While no lines were identified in the data cubes, continuum emission at 1.2 mm was detected from four sources at \( 5\,\sigma \) significance levels. Among the four ALMA sources, only HPC1001.a is slightly resolved, while the rest are point-like. We thus adopted aperture photometry for HPC1001.a and peak fluxes for HPC1001.b, d, and e.

HPC1001 was also detected (but barely resolved) in Herschel, SCUBA2 (Simpson et al., 2019), and MeerKAT (Jarvis et al. 2016) imaging (Fig. 3-left). We measured its integrated far-IR (FIR) to radio photometry by performing the superdeblending technique as described in Jin et al. (2018) at the fixed position of HPC1001.a. The Herschel colors (\( S_{250 \mu m} < S_{350 \mu m} < S_{500 \mu m} \)) indicate that it is a 500 \( \mu m \) riser which is even redder than the cluster CL J1001 at \( z = 2.5 \), supporting its \( z > 2 \) nature (Riechers et al. 2017; Donevski et al. 2018; Cairms et al. 2022).

In order to constrain the IR star formation rate (SFR) and dust mass (\( M_{\text{dust}} \)) of the group, we fit the integrated FIR to radio photometry using the STARDUST code (Kokorev et al. 2021). The integrated SED (Fig. 3-right) was well fitted by a starburst template (GN20, Magdis et al. 2012) at \( z = 3.7 \) with dust-obscured SFR\(_{\text{IR}} = 743 \, M_\odot \) yr\(^{-1}\). Assuming that all ALMA detected sources share the same SED shape as the integrated SED, we derived SFRs and dust masses for the four ALMA sources by splitting the total values based on the ALMA 1.2 mm flux of each source (Fig. 3-right and Table 1). Given the ALMA brightness and NIR faintness, HPC1001.a could be
**Fig. 1.** Galaxy overdensity HPC1001 at $z \approx 3.7$. Left: VISTA $K_s$ image overlaid with SCUBA2 850$\mu$m contours in 3–5$\sigma$ levels. Galaxies at $3.5 < z < 3.9$ are marked with red circles and labeled with phot-z. The dashed circle shows the virial radius of a dark matter halo $M_{\text{DM}} = 10^{13} M_\odot$. Right: HSC $i$ and VISTA $JK_s$ color image for the core region, overlaid by ALMA 1.2 mm continuum contours in yellow (3–5$\sigma$). The ALMA beam ($0.53'' \times 0.44''$, PA = $-25.8^\circ$) is shown with a yellow ellipse in the bottom left. Candidate members are labeled with text in orange.

Galaxy QSO #b, z=3.56 #c, z=3.67 #d, z=3.54 #e, z=3.67 #g, z=3.68 #f, z=3.86 #i, z=3.56 #j, z=3.79

**Fig. 2.** Optical and NIR SEDs from the COSMOS2020 catalog. Each SED was fit by both a galaxy template (red curve) and a QSO template (yellow curve). Lines on subpanels mark the Bayesian (red) and $\chi^2$ (blue) PDF($z$) of the galaxy template fitting, respectively. The best fit phot-z estimates are also quoted in each panel.

3. Results

3.1. Dark matter halo mass

We estimated the dark matter halo mass $M_{\text{DM}}$ of HPC1001 with six methods: (1) Using the $M_{\text{halo}}$-$M_*$ scaling relation from a starburst with optically thick dust emission, in which case the inferred dust mass is likely to be overestimated by a factor of two (Cortzen et al. 2020; Jin et al. 2022). This consideration has been incorporated in the adopted dust mass uncertainty of the source (Table 1).
Table 1. Physical properties of HPC1001.

| Name     | ID (a) | $z_{\text{phot,NIR}}$ | $S_{1.2\,\text{mm}}$ [mJy] | log($M_*/M_\odot$) | SFR$_{\text{UV,cov}}$ [M_\odot yr$^{-1}$] | SFR$_{\text{IR}}$ [M_\odot yr$^{-1}$] | log($M_{\text{dust}}/M_\odot$) | $f_{\text{gas}}$ [M_\odot/M_\odot] |
|----------|--------|------------------------|-----------------------------|---------------------|------------------------------------------|--------------------------------|----------------------------------|----------------------------------|
| HPC1001.a | –      | 3.7 (b)                | 1.42 ± 0.06                 | 10.3 ± 0.2          | –                                        | 429 ± 20                      | 8.9 ± 0.05                       | 4.3 ± 2.4                       |
| HPC1001.b | 54765  | 3.56 ± 0.08            | 0.35 ± 0.06                 | 11.0 ± 0.1          | 110$^{+25}_{-6}$                         | 106 ± 20                      | 8.32 ± 0.08                      | 0.2 ± 0.1                       |
| HPC1001.c | 54989  | 3.67 ± 0.07            | <0.18                       | 9.8 ± 0.1           | 17$^{+3}_{-8}$                           | <59                          | <8.07                           | <1.9                            |
| HPC1001.d | 55311  | 3.54 ± 0.13            | 0.32 ± 0.06                 | 10.5 ± 0.1          | 64$^{+12}_{-11}$                         | 97 ± 20                       | 8.28 ± 0.09                      | 0.6 ± 0.2                       |
| HPC1001.e | 55655  | 3.67 ± 0.17            | <0.18                       | 10.2 ± 0.1          | 50$^{+46}_{-35}$                         | 115 ± 20                      | 8.36 ± 0.09                      | 1.4 ± 0.5                       |
| HPC1001.f | 55750  | 3.86 ± 0.10            | <0.19                       | 10.2 ± 0.2          | 32$^{+29}_{-32}$                         | <61                          | <8.07                           | <0.7                            |
| HPC1001.g | 56086  | 3.68 ± 0.06            | <0.19                       | 9.5 ± 0.1           | 13$^{+2}_{-3}$                           | <61                          | <8.07                           | <3.7                            |
| HPC1001.h | 880059 (c) | 3.79 ± 0.37          | <0.18                       | 10.0 ± 0.2          | 77$^{+42}_{-35}$                         | <59                          | <8.07                           | <0.9                            |
| HPC1001.i | 32841  | 3.56 ± 0.29            | <0.34                       | 10.2 ± 0.2          | 29$^{+16}_{-2}$                          | <106                         | <8.33                           | <1.3                            |
| HPC1001.j | 74590  | 3.79 ± 0.12            | <0.33                       | 10.6 ± 0.1          | 73$^{+15}_{-37}$                         | <126                         | <8.31                           | <0.5                            |
| Total    | HPC1001| –                      | 3.65 ± 0.1                  | 2.46 ± 0.05         | 11.4 ± 0.2                               | 391$^{+66}_{-47}$            | 743 ± 3                          | 9.17 ± 0.05                      |

Notes. $S_{1.2\,\text{mm}}$: ALMA 1.2 mm photometry. SFR$_{\text{UV,cov}}$: UV-corrected SFR from the COSMOS catalogs. (a) COSMOS2020 Farmer catalog ID. (b) We adopted $z = 3.7$ from the weighted average of the group. (c) This source is not in the COSMOS2020 catalog, and we list parameters from the COSMOS2015 catalog instead. (d) We did not account for HPC1001.h due to its inconsistency between SFR$_{\text{UV,cov}}$ and SFR$_{\text{IR}}$.

Behroozi et al. (2013) and the stellar mass of the most massive source HPC1001.b, it yields a lower limit for the halo mass of log($M_{\text{DM}}/M_\odot$) $> 13$; (2) following the methodology presented in Daddi et al. (2021, 2022), we estimated the halo mass based on the stellar masses above the completeness limit of the COSMOS survey, that is log($M_*/M_\odot$) $> 9.7$ at $z = 3.7$ (Weaver et al. 2022). This mass limit holds for nine out of ten members, and the sum of stellar masses is $M_* = 2.45 \times 10^{11} M_\odot$. We then extrapolated a total stellar mass of $M_{\text{stellar}} = 2.6 \times 10^{11} M_\odot$ down to $10^7 M_\odot$ assuming the stellar mass function of field galaxies at $3 < z < 4$ (Muzzin et al. 2013). Adopting the dynamical mass-constrained $M_{\text{halo}} - M_*$ scaling relation for $z \sim 1$ clusters with $0.6 \times 10^{14} < M_{\text{halo}}/M_\odot < 16 \times 10^{14}$ (van der Burg et al. 2014) yields a halo mass of log($M_{\text{halo}}/M_\odot$) $= 12.8$; (3) adopting the stellar-to-halo mass relation of Shuvtov et al. (2022) and $M_{\text{stellar}} = 2.6 \times 10^{11} M_\odot$, we obtained a halo mass of log($M_{\text{DM}}/M_\odot$) $= 12.7$; (4) we used the background and point source subtracted combined Chandra+XMM-Newton image in the 0.5–2 keV (Gozalıasl et al. 2019) to study the X-ray emission of HPC1001. No extended source was detected at the position of the group. Using a 32″ radius aperture, we placed two sigma upper limit $S_{0.5–2\,\text{keV}} < 4.2 \times 10^{-16}$ ergs s$^{-1}$ cm$^{-2}$, which yielded an upper limit of $M_{\text{200}} < 2.9 \times 10^{13} M_\odot$; (5) assuming a group velocity dispersion $\sigma_V = 400$ km s$^{-1}$, we found that the galaxy number density of HPC1001 (in putative $R_{\text{vir}} < 20$′′) is more overdense than the average by a factor of 1.80 at $z = 3.7$ in COSMOS2020. Applying a mean baryon and dark matter density of 7.77 $\times 10^{-26}$ kg m$^{-3}$ in comoving volume from Planck cosmology and a galaxy bias factor of 10–20 at $z = 3.7$ (Tinker et al. 2010), we obtained a halo mass of log($M_{\text{DM}}/M_\odot$) $= 13–13.4$; (6) assuming that HPC1001 is the most extreme halo in COSMOS at $z > 3.6$, that is to say that it is the only case that can be found in a 2 deg$^2$ field, the halo mass would be $M_{\text{DM}} \sim 2.7 \times 10^{13} M_\odot$ in Planck cosmology.
Based on these estimates, we adopted an average halo mass of $\log(M_{\text{DM}}/M_\odot) = 13$ with an uncertainty of 0.4 dex that is representative at these faint levels (e.g., Looser et al. 2021; Daddi et al. 2022). Assuming the structure is collapsed, the virial radius corresponding to this halo mass would be 140 kpc (Goerdt et al. 2010), which is large enough to enclose all ten sources that are included in deriving the halo mass.

In Fig. 4-left, we compared the halo mass of HPC1001 to literature data taken from Daddi et al. (2022). With a halo mass that is lower than that of massive clusters, such as CL J1001, but comparable to that of galaxy groups, such as RO-0958 at $z = 3.3$ and RO-0959 at $z = 3.1$ (Daddi et al. 2022), we propose that HPC1001 is likely to be a galaxy group. Its halo mass is above the mass criterion for generating shocks, and it is well within the region where cold streams are accreting in hot media, as predicted by simulations (Dekel et al. 2013). In this case, cold gas inflow is taking place in HPC1001 and similar to RO-1001 and RO-0959 (Daddi et al. 2021, 2022), this inflow should be detectable via diffuse Lyα emission.

### 3.2. Star formation activity and dust fraction

In order to obtain the total SFR within the HPC1001 structure, we summed the SFR$_{\text{HP}}$ of the four ALMA-detected sources with the SFR$_{\text{UV,corr}}$ of the remaining unobscured sources. The total UV+FIR star formation rate is SFR$_{\text{total}} \approx 900 M_\odot$ yr$^{-1}$. Applying the scaling relations of Daddi et al. (2022), for a dark matter halo with $\log(M_{\text{DM}}/M_\odot) = 13$ at $z = 3.7$, the expected cold baryonic accretion rate (BAR) is $2780 M_\odot$ yr$^{-1}$ which in turn corresponds to a SFR of $800 M_\odot$ yr$^{-1}$. The total SFR of HPC1001 agrees with the model prediction, further supporting the scenario that HPC1001 is a cold-accretion-fed group.

Moving the focus to individual sources, we find that HPC1001.b has a specific star formation rate (sSFR) that places the source $\times3.0 \pm 1.0$ below the main sequence at the corresponding redshift (Fig. 4-middle). At the same time, as shown in Fig. 4-right, its dust to stellar mass ratio $f_{\text{dust}}$ is lower by a factor of $\times4-7$ with respect to that of main sequence galaxies at $z = 3.7$ (Tacconi et al. 2020; Kokorev et al. 2021) and comparable to that of $z = 1.5$ quiescent galaxies (Gobat et al. 2018; Magdis et al. 2021; Caliendo et al. 2021). Assuming $M_{\text{gas}} = 100 \times M_{\text{dust}}$ and accounting for a 20% uncertainty due to unknown metallicity (Jin et al. 2019, 2022), HPC1001.b would have an inferred gas fraction of $M_{\text{gas}}/M_\star = 0.2 \pm 0.1$ and a gas depletion time of 200 Myr.

### 4. Discussion

#### 4.1. A galaxy group in maturing phase at $z \approx 3.7$

Multiple pieces of evidence support the idea that HPC1001 is a galaxy group at $z \approx 3.7$: (1) the high sky overdensity of galaxies; (2) the excellent agreement of photometric redshifts from four independent methods; and (3) the fact that the inferred dark halo mass and SFR are comparable with that of galaxy groups at $z > 3$ (Daddi et al. 2022). We note that the large structure surrounding HPC1001 hosts extra NIR-detected galaxies at a similar redshift ($3.3 < z < 4$) that are coupled with an overdensity of submillimeter sources (Fig. A.2). These sources could be in sub-halos that associate with HPC1001 in a megaparsec-scale structure (Koyama et al. 2013; Cucciati et al. 2018; Jin et al. 2021), which needs to be identified by future spectroscopic follow-ups.

On the other hand, the remarkably low gas and dust fractions of HPC1001.b suggest that this galaxy is more evolved compared to normal star-forming galaxies at the same epoch (Gómez-Guijarro et al. 2019). These properties are in line with the maturing phase of protoclusters (Shimakawa et al. 2018), where the core hosts both starbursting and quenching members, the overall star formation activity begins to decline, and the dark matter halo starts to collapse. This scenario has not been observed at $z > 3$, and if confirmed HPC1001 would be the earliest structure in maturing phase detected to date.

#### 4.2. A candidate for a galaxy in quenching

HPC1001.b is the most massive galaxy in the structure, with a stellar mass comparable to that of the quiescent cluster members at $z \sim 3$ (Kubo et al. 2021; Kalita et al. 2021). Assuming no gas inflow, HPC1001.b would deplete its molecular gas reservoir in 200 Myr and become quiescent at $z > 3.2$ with $M_\star \sim 1.2 \times 10^{11} M_\odot$, forming the first generation of massive quiescent galaxies in dense environments. Of course, potential cold gas inflow (Fig. 4-left) can replenish the star-forming gas reservoir and maintain star formation in HPC1001.b. However, the presence of a quenched galaxy in the RO-1001 group where gas inflow is present (Kalita et al. 2021) suggests that quiescent...
galaxies can form in high-$z$ dense environments despite the ongoing gas inflow.

In order to fully characterize the nature of HPC1001 and explore the quiescence of HPC1001.b, [CII]158 µm observations are crucial to infer spectroscopic redshifts, confirm group members, and robustly measure gas masses and SFRs. The kinematic information from [CII] or [CI] can also verify if HPC1001 is compatible with a single virialized structure. On the other hand, VLT/MUSE and GTC/MEGARA can be employed to detect the cold gas inflow via Ly$\alpha$ observations. Unfortunately, HPC1001 is not covered by the JWST COSMOS-Web survey. However, a future follow-up observation of the structure with JWST/NIRCam and NIRSpec will be a powerful tool to reveal potential inter-cluster light and measure the metallicity of the group members.

Acknowledgements. This paper makes use of the ALMA data: ADS/JAO.ALMA 2013.1.00034.S. ALMA is a partnership of ESO (representing its member states), NSF (USA), and NINS (Japan), together with NRC (Canada), MOST and ASIAA (Taiwan), and KASI (Republic of Korea), in cooperation with the Republic of Chile. The Joint ALMA Observatory is operated by ESO, AUI/NAOJ, and NRAO. The Cosmic Dawn Center (DAWN) is funded by the Danish National Research Foundation under grant No. 140. SJ is supported by the European Union’s Horizon research and innovation program under the Marie Skłodowska-Curie grant agreement No. 101060888. SJ and GEM acknowledge financial support from the Villum Young Investigator grant 37440 and 13160. TRG acknowledges support from the Carlsberg Foundation (grant no CF20-0534).

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Acknowledgements. This paper makes use of the ALMA data: ADS/JAO.ALMA 2013.1.00034.S. ALMA is a partnership of ESO (representing its member states), NSF (USA), and NINS (Japan), together with NRC (Canada), MOST and ASIAA (Taiwan), and KASI (Republic of Korea), in cooperation with the Republic of Chile. The Joint ALMA Observatory is operated by ESO, AUI/NAOJ, and NRAO. The Cosmic Dawn Center (DAWN) is funded by the Danish National Research Foundation under grant No. 140. SJ is supported by the European Union’s Horizon research and innovation program under the Marie Skłodowska-Curie grant agreement No. 101060888. SJ and GEM acknowledge financial support from the Villum Young Investigator grant 37440 and 13160. TRG acknowledges support from the Carlsberg Foundation (grant no CF20-0534).

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Appendix A: Galaxy over-density

![Galaxy Sky Density](image)

**Fig. A.1.** Galaxy sky density of (proto)cluster cores in comoving area. Densities were measured in a $10'\times10'$ area with the maximum number of cluster members covered by the area, where we applied a stellar mass limit $\log(M_*/M_{\odot}) > 10.2$ for optical-detected members and a flux limit $S_{1850} > 0.3$ mJy for 1 mm continuum detections, respectively. We show extreme cases in literature from Wang et al. (2016), Miller et al. (2018), Willis et al. (2020), and Daddi et al. (2021). We note that CL J1001 and DRC lack 1 mm observations; their 1 mm fluxes were derived by scaling their 870$\mu$m and 2 mm continuum under assumption of a GN20 dust template (Magdis et al. 2012). HPC1001 is highlighted in green, showing the highest density at $z > 3$.

Appendix B: Redshift comparison

![Redshift Comparison](image)

**Fig. B.1.** Comparison of photo-z between the four versions in the COSMOS2020 catalog (Weaver et al. 2022) and the COSMOS2015 catalog (Laigle et al. 2016). The weighted averages are $z = 3.65 \pm 0.07$ (Farmer LePhare), $z = 3.68 \pm 0.02$ (Farmer EAZY), $z = 3.63 \pm 0.08$ (Classic LePhare), $z = 3.62 \pm 0.04$ (Classic EAZY), and $z = 3.72 \pm 0.08$ (COSMOS2015), respectively.

![SCUBA2 S/N Map](image)

**Fig. A.2.** SCUBA2 850$\mu$m S/N map. The rms is shown with text and the S/N is indicated by the color bar from 2.5 to 7$\sigma$. Green circles mark galaxies that have photo-z $3.3 < z < 4$ and match SCUBA2 detections ($S/N_{850} > 3$ and tolerance $< 10''$). HPC1001 is in an overdensity of sub-millimeter galaxies at a similar redshift.