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The NuSTAR Serendipitous Survey: Hunting for the Most Extreme Obscured AGN at >10 keV

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Abstract

We identify sources with extremely hard X-ray spectra (i.e., with photon indices of $\Gamma \lesssim 0.6$) in the 13 deg$^2$ NuSTAR serendipitous survey, to search for the most highly obscured active galactic nuclei (AGNs) detected at >10 keV. Eight extreme NuSTAR sources are identified, and we use the NuSTAR data in combination with lower-energy X-ray observations (from Chandra, Swift XRT, and XMM-Newton) to characterize the broadband (0.5–24 keV) X-ray spectra. We find that all of the extreme sources are highly obscured AGNs, including three robust Compton-thick (CT; $N_{\text{H}} > 1.5 \times 10^{22}$ cm$^{-2}$) AGNs at low redshift ($z < 0.1$) and a likely CT AGN at higher redshift ($z = 0.16$). Most of the extreme sources would not have been identified as highly obscured based on the low-energy (<10 keV) X-ray coverage alone. The multiwavelength properties (e.g., optical spectra and X-ray–mid-IR luminosity ratios) provide further support for the eight sources being significantly obscured. Correcting for absorption, the intrinsic rest-frame 10–40 keV luminosities of the extreme sources cover a broad range, from $\approx 5 \times 10^{42}$ to $10^{45}$ erg s$^{-1}$. The estimated number counts of CT AGNs in the NuSTAR serendipitous survey are in broad agreement with model expectations based on previous X-ray surveys, except for the lowest redshifts ($z < 0.07$), where we measure a high CT fraction of $f_{\text{CT}}^{\text{obs}} \approx 30^{+16}_{-12}$$\%$. For the small sample of CT AGNs, we find a high fraction of galaxy major mergers (50% ± 33%) compared to control samples of “normal” AGNs.

Key words: galaxies: active – galaxies: nuclei – quasars: general – surveys – X-rays: galaxies

1. Introduction

The majority of cosmic supermassive black hole growth has occurred in an obscured phase (e.g., see Brandt & Alexander 2015, for a review), during which gas and dust cover the central active galactic nucleus (AGN). Historically, the importance of highly obscured AGNs has been inferred from the shape of the extragalactic cosmic X-ray background (CXB), the high-energy hump of which (peaking at $\approx 20–30$ keV) requires significant populations of either highly obscured or reflection-dominated systems (e.g., Setti & Woltjer 1979; Comastri et al. 1995; Gilli et al. 2007; Treister et al. 2009). Large population studies have now quantified the relative abundance of obscured and unobscured black hole growth phases (e.g., Aird et al. 2015; Buchner et al. 2015).
substantial fraction of the growth appears to occur during the most obscured “Compton-thick” (CT) phases, where the absorbing column density exceeds the inverse of the Thomson scattering cross section \( (N_H \gtrsim 1.5 \times 10^{24} \text{ cm}^{-2}) \). However, the intrinsic absorption distribution of AGNs has proven difficult to constrain, especially at the highly obscured to CT end, where AGNs are particularly challenging to identify.

Besides completing a census, identifying the most highly obscured AGNs is crucial to our understanding of the environment of supermassive black hole growth. The unified model of AGNs (e.g., Antonucci 1993; Urry & Padovani 1995; Netzer 2015), which largely succeeds at describing AGNs in the local universe, posits that unobscured, obscured, and CT systems have intrinsically similar nuclear structures but are simply viewed from different inclination angles. In tension with this model (at least in its simplest form) are observational results that find possible evidence for high merger fractions in highly obscured AGN samples (e.g., Kocevski et al. 2015; Del Moro et al. 2016; Koss et al. 2016a; Ricci et al. 2017). Furthermore, observations of the clustering of AGNs find that obscured and unobscured AGNs may inhabit different large-scale environments (e.g., Allevato et al. 2011, 2014; DiPompeo et al. 2014, 2016; Donoso et al. 2014; but see also Mendez et al. 2016; Ballantyne 2017). These results may suggest that high AGN obscuration can be linked to specific phases in the galaxy–AGN coevolutionary sequence (e.g., Sanders et al. 1988; Hopkins et al. 2008; Alexander & Hickox 2012), potentially associated with periods of rapid black hole growth (e.g., Draper & Ballantyne 2010; Treister et al. 2010).

A challenge in answering these questions is that most wavelength regimes are subject to strong biases against detecting highly obscured AGNs, due to a combination of (i) line-of-sight extinction and (ii) dilution by light from other (e.g., stellar) processes. Selection methods exist that are relatively unhindered by (i), such as mid-infrared (MIR) color selection (e.g., Lacy et al. 2004; Stern et al. 2005; Daddi et al. 2007; Fiore et al. 2008; Mateos et al. 2012; Stern et al. 2012) and optical spectroscopic selection based on high-ionization emission lines (e.g., Zakamska et al. 2003; Reyes et al. 2008). However, these techniques both suffer from (ii), especially at sub-quasar luminosities, and both still require X-ray follow-up of the AGNs to provide accurate measurements of the line-of-sight gas column densities (e.g., Vignali et al. 2006; Civano et al. 2007; Alexander et al. 2008; Vignali et al. 2010; Jia et al. 2013; LaMassa et al. 2014; Del Moro et al. 2016). Hard (>10 keV) X-ray observations, on the other hand, have the advantage of very little dilution from other processes and are relatively unaffected by line-of-sight obscuring material up to CT levels of absorption.

For high redshift AGNs \((z \gtrsim 2)\) soft X-ray telescopes (e.g., Chandra and XMM-Newton) sample the rest-frame hard X-ray band and are therefore effective in identifying the features of CT absorption (e.g., Comastri et al. 2011; Brightman et al. 2014). In the lower-redshift universe, however, hard X-ray telescopes become essential. Large (e.g., all-sky) surveys with nonfocusing hard X-ray missions (e.g., Swift BAT and INTEGRAL) have been important for the identification of highly obscured AGNs in the very local universe \((z < 0.05\) e.g., Burlon et al. 2011; Vasudevan et al. 2013; Ricci et al. 2015; Akylas et al. 2016; Koss et al. 2016a). Now, with the first focusing hard X-ray mission (NuSTAR; Harrison et al. 2013) it is possible to study source populations that are approximately two orders of magnitude fainter, thus extending to lower luminosities and higher redshifts. The largest extragalactic survey being undertaken with NuSTAR is the serendipitous survey (Alexander et al. 2013; Lansbury et al. 2017), which has covered \(\approx 13\) deg\(^2\) and detected 497 sources, 276 of which have spectroscopic redshifts. The areal coverage and sample size are large compared to the dedicated NuSTAR extragalactic blankfield surveys (e.g., in the ECDFS and COSMOS fields; Civano et al. 2015; Mullaney et al. 2015), making the serendipitous survey well suited to the discovery of rare populations such as CT AGNs. The latter have proven elusive in the NuSTAR surveys thus far, with only one to two high-confidence CT AGNs being identified overall (e.g., Civano et al. 2015; Del Moro et al. 2017; Zappacosta et al. 2017).

In this paper, we search for the most extreme hard X-ray sources in the 40-month NuSTAR serendipitous survey sample, and as a result we reveal new robust CT AGNs. First, we select the objects with the highest NuSTAR band ratios, implying very hard spectral shapes and hence the likely presence of heavy absorption. Although band ratios only give a crude estimate of absorption, they are nevertheless an effective way to isolate the most extreme outliers (e.g., Koss et al. 2016a). Second, we perform a detailed analysis of the X-ray and multwavelength properties of these extreme objects and discuss how their properties compare to those of the general AGN population. The paper is structured as follows. Section 2 describes the selection of the eight extreme objects from the NuSTAR serendipitous survey sample. Section 3 details the data used and the soft X-ray counterparts. In Section 4 we characterize the X-ray spectra of the sources (Section 4.1) and present the results for the X-ray spectral properties (Section 4.2). In Section 5 we investigate potential independent estimates of the source obscuration properties through indirect techniques. Section 6 presents the optical properties of the sample, including a summary of the optical spectral properties (Section 6.1) and host galaxy imaging, with a focus on the frequency of galaxy mergers (Section 6.2). In Section 7 we discuss the CT AGNs and their implications for the prevalence of CT absorption within the broader hard-X-ray-selected AGN population. Finally, our main results are summarized in Section 8. The cosmology adopted is \((\Omega_M, \Omega_{\Lambda}, h) = (0.27, 0.73, 0.70)\). All uncertainties and limits are quoted at the 90% confidence level (CL), unless otherwise stated.

2. The Sample of Extreme, Candidate Highly Obscured AGNs from the NuSTAR Serendipitous Survey

We start with the total 40-month NuSTAR serendipitous survey sample (497 sources; Lansbury et al. 2017). To select sources with extremely hard X-ray spectra compared to the rest of the NuSTAR serendipitous survey sample, we identify sources with high hard-to-soft band ratios (BR\(_{\text{Nu}}\)), calculated as the ratio of the 8–24 keV to 3–8 keV count rates. We apply a cut at \(\text{BR}_{\text{Nu}} > 1.7\) (see Figure 1), which corresponds to an effective (i.e., observed) photon index of \(\Gamma_{\text{eff}} \lesssim 0.6\). This cut is motivated by the \(\text{BR}_{\text{Nu}}\) values observed for CT AGNs in other NuSTAR programs (e.g., Baloković et al. 2014; 28). The power-law photon index \((\Gamma)\) is defined as follows: \(F_E \propto E^{-\Gamma}\), where \(F_E\) is the photon flux and \(E\) is the photon energy.
Here we comment on the maximum energies at which the sources are detected with NuSTAR. Table 1 lists the standard NuSTAR energy bands (i.e., the full, soft, and hard bands) for which sources are detected. By selection, all eight sources are detected in the 8–24 keV band. Splitting this hard band into sub-bands of 8–16 keV and 16–24 keV, all eight sources are detected in the former band, and all except J1444 and J1653 are detected in the latter band. For the six sources detected at 16–24 keV, the highest and lowest Poisson false probabilities are $P_{\text{False}} = 2 \times 10^{-3}$ and $10^{-8}$, respectively (i.e., the detections range from $\approx 3\sigma$ to highly significant). Only one source shows evidence for emission at $>24$ keV: J1506, which is detected in the 24–50 keV band at the $\approx 3\sigma$ significance level. In summary, two sources are detected up to a maximum energy of $\approx 16$ keV, five sources are detected up to $\approx 24$ keV, and a single source is weakly detected at even higher energies (up to $\approx 50$ keV).

### 2.1. A Note on Associated Sources

Six out of eight (75%) of the sources in this sample were serendipitously detected in NuSTAR observations of bright low-redshift Swift BAT AGNs. The three serendipitous NuSTAR sources J0505, J1506, and J1512 are likely to be weakly associated with the brighter BAT AGNs: each source lies within $\pm 500$ km s$^{-1}$ of the redshift of the BAT AGN and at a projected separation of $\approx 150$–550 kpc. The associations are “weak” in that the physical separations are large, and the sources are therefore unlikely to be interacting. The associated redshifts are unlikely to occur by chance given that hard X-ray sources at these flux levels ($f_{\text{24 keV}} = 2 - 6 \times 10^{-13}$ erg s$^{-1}$ cm$^{-2}$), and within $\pm 500$ km s$^{-1}$ of the BAT redshifts, have very low sky densities of $\approx 0.01$ deg$^{-2}$ (e.g., Treister et al. 2009). The latter implies low chance coincidence rates of $\approx 10^{-3.5}$ for the three cases of associated redshifts above. The effect of these weak associations on number counts measurements for CT AGNs is accounted for in Section 7.

In the overall 40-month NuSTAR serendipitous survey, redshift associations between serendipitous sources and science targets like the above are rare (Lansbury et al. 2017). The exception is at $z < 0.07$, where 10 out of 15 sources (including J0505, J1506, and J1512) show evidence for associations. We emphasize, however, that the majority of extragalactic sources in the NuSTAR serendipitous survey (247/262 of the spectroscopically identified sample) are at higher redshifts ($z > 0.07$), meaning that number counts measurements for the survey (e.g., Harrison et al. 2016) are not impacted.

### 3. Data

Table 2 provides details of the NuSTAR and soft (<10 keV) X-ray (i.e., Chandra, Swift XRT, and XMM-Newton) data sets used in this work. For each source we adopt the soft X-ray observatory data that provide the most sensitive coverage at $<10$ keV. For four sources this coverage is from Swift XRT, for three sources it is from XMM-Newton, and for one source it is from Chandra. For five sources we use the combined soft X-ray data set from multiple individual observations (as

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29 Sources are classed as associated if their velocity offset from the science target $[\Delta v(z)]$ is smaller than 5% of the total science target velocity (see Lansbury et al. 2017).

50 At $z > 0.07$ only two sources are flagged as associated.
detailed in Table 2) to obtain the most precise X-ray constraints possible. The soft X-ray observations are generally not contemporaneous with the NuSTAR observations. Section 4.1 discusses the possibility of variability for these sources.

3.1. Soft X-Ray Counterparts to the Extreme NuSTAR Sources

The soft X-ray counterparts improve the X-ray positional accuracy and, when combined with the NuSTAR data, allow for accurate spectral constraints using the broadest energy band possible. Of the eight extreme NuSTAR sources studied here, two lack soft X-ray counterparts (J1410 and J1506). In these cases there is no Chandra or XMM-Newton coverage, and the sources are undetected in the combined archival Swift XRT coverage (running wavdetect with a detection threshold of 10^−5). The other six extreme NuSTAR sources have identified soft X-ray counterparts. For five of these (J0505, J0823, J1444, J1512, and J1653) the soft X-ray counterparts are identified in Lansbury et al. (2017). Since J0505 has two XMM-Newton sources nearby to the NuSTAR source, we provide evidence below to support our correct counterpart identification in this case. For the remaining source (J1534), the Chandra counterpart is faint and did not satisfy the detection criteria in Lansbury et al. (2017). Below we detail the identification of this specific counterpart.

For J0505, there are two potential counterparts in the 3XMM catalog, one at 14″ offset from the NuSTAR position (R.A. = 76°49983, decl. = −23°83536; hereafter “XMM1”) and one brighter source at 27″ offset (R.A. = 76°49296 decl. = −23°82597; hereafter “XMM2”). To examine the X-ray spectra, we used the MOS data for XMM1 (the source lies on a chip gap for PN) and the PN plus MOS data for XMM2. The 0.5–10 keV spectrum for XMM1 is extremely flat (with an effective photon index of Γeff = −0.9±0.5), and there is a line detection consistent with Fe Kα (rest-frame E = 6.3 ± 0.1 keV). The Fe Kα line has a high equivalent width of EWFeKα = 1.4±0.9 keV, suggesting a highly absorbed AGN. For XMM2, the 0.5–10 keV spectrum is steeper (Γeff = 1.4 ± 0.2). Although XMM2 is brighter than XMM1 over the full energy band, XMM1 is significantly brighter for the energies at which NuSTAR is sensitive: for the 3–10 keV energy band, XMM1 and XMM2 have fluxes of 8.9 × 10^−14 erg s^−1 cm^−2 and 1.8 × 10^−14 erg s^−1 cm^−2, respectively. Given these fluxes and the relative spectral slopes of XMM1 and XMM2 (with the former sharply increasing, and the latter decreasing, toward higher X-ray energies), and the fact that the majority of NuSTAR source counts (79%) lie at high energies (>8 keV), we expect XMM1 to dominate the NuSTAR-detected emission. We therefore adopt XMM1 as the counterpart to J0505.

For J1534, the deepest soft X-ray coverage is from a 171.5 ks Chandra observation (obsID 16092, which targeted Arp 220). Running wavdetect for the broad Chandra energy band of 0.5–7 keV, no sources are blindly detected within the NuSTAR error circle with false probabilities (i.e., sigmethresh values) of P_{false} ≪ 10^−4. However, running the source detection for multiple energy bands, there is a significant detection at 0.5–2 keV, with P_{false} ≈ 10^−6. Adding further confidence to the reliability of this source, Sloan Digital Sky Survey (SDSS) coverage reveals a prominent z = 0.160 galaxy within the NuSTAR error circle (SDSS J153445.80+233121.2), which agrees with the Chandra position within the positional uncertainties (0″.6 offset). For an independent assessment of the significance of the Chandra source, we perform aperture

![Figure 2. NuSTAR and soft X-ray (Chandra, Swift XRT, and XMM-Newton) images for the eight extreme NuSTAR serendipitous survey sources. Each column corresponds to an individual NuSTAR source (the abbreviated source names are shown). 30″-radius circular apertures are shown for each source, centered on the NuSTAR position. Upper two rows: NuSTAR hard-band (8–24 keV) images, both smoothed (with a top hat function of radius 14 pixels; first row) and unsmoothed (second row). Lower two rows: soft X-ray images from Chandra (the 0.5–2 keV band is shown for J1534), XMM-Newton (the full energy band is shown for J0505, J0823, and J1653), and Swift XRT (the full energy band is shown for J1410, J1444, J1506, and J1512). The data are shown both smoothed (with a Gaussian function of radius 3 pixels; third row) and unsmoothed (fourth row). The soft X-ray counterpart positions are marked by smaller (10″ radius) circular apertures, for all of the sources except J1410 and J1506 (which are undetected in the available Swift XRT coverage; see Section 4).]
photometry (2″ source radius; large background annulus) at the SDSS position. For the 0.5–2 keV band, the source is indeed detected at the 4σ level (according to the binomial false probability). The NuSTAR/Chandra flux ratio for J1534 is extremely high (e.g., \( f_{2-4} / f_{0.5-2} = 141 \)). For comparison, four sources in the NuSTAR-COSMOS survey have similarly high flux ratios (ranging from \( f_{2-4} / f_{0.5-2} = 100 \) to 304), and all of these have been identified as highly obscured AGNs (e.g., Brightman et al. 2014; Lanzuisi et al. 2015; Zappacosta et al., 2017). The Chandra spectrum for J1534 is further discussed in Section 4.1.

### 3.2. X-Ray Spectroscopic Products

The NuSTARDAS task nuproducts is used to extract NuSTAR source spectra, background spectra, and response files. We adopt circular source extraction regions of 45″ radius where possible, and of 30″ radius for two cases where the source is either close to a bright science target or to the field-of-view (FOV) edge. We perform separate spectral extractions for the two individual NuSTAR telescopes (FPMA and FPMB). For J0823, we limit the modeling to FPMB, since the source is only fully within the NuSTAR FOV for FPMB.

For the six sources with soft X-ray counterparts, we extract additional spectra from the archival soft X-ray data sets detailed in Table 2, using the relevant software packages (the Chandra Interactive Analysis Observations software, \(^3\)) the Swift XRT analysis software distributed with HEASoft, \(^3\) and the XMM-Newton Science Analysis Software \(^4\). We adopt source extraction apertures of 5″, 10″, and 12″–15″ radius for the Chandra, Swift XRT, and XMM-Newton data, respectively. For J1444 we co-add the Swift XRT spectra across all six observations, since the source is only significantly detected in the co-added data.

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### Table 1

| Field Name | R.A. (3) | Decl. (4) | \( z \) (5) | BR\(_{50} \) (6) | Det. (7) | Nu\(_{\text{flat}} \) (8) | Source Type (9) |
|------------|----------|----------|-------------|--------------|----------|----------------|-----------------|
| NuSTAR J050559-2349.9 | J0505 | 76.49839 | −23.83169 | 0.036 | >3.8 | F H | 0.2 | 2MASX J05054575-235113 |
| NuSTAR J082303-0502.7 | J0823 | 125.76385 | −5.04650 | 0.313 | >2.0 | F H | 0.5 | FAIRALL 0272 |
| NuSTAR J141056-4230.0 | J1410 | 212.73727 | −42.50139 | 0.067 | 1.9 ± 0.8 | F S H | 0.5 | 2MASX J14104482-422832 |
| NuSTAR J144406+2506.3 | J1444 | 221.0820 | 25.10515 | 1.539 | >2.3 | F H | 0.3 | PKS 1441+25 |
| NuSTAR J150645+0346.2 | J1506 | 236.6904 | 3.77118 | 0.034 | >4.2 | F H | 0.4 | 2MASX J15064412+035144 |
| NuSTAR J151253-8124.3 | J1512 | 228.22497 | −81.40501 | 0.069 | 1.8 ± 0.6 | F S H | 1.0 | 2MASX J15144217-812337 |
| NuSTAR J153445+2331.5 | J1534 | 353.67836 | 23.52593 | 0.160 | >3.5 | H | 0.4 | Arp 220 |
| NuSTAR J165346+3953.7 | J1653 | 253.4413 | 39.89639 | 0.354 | >2.7 | H | 0.2 | Mrk 501 |

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#### 4. X-Ray Properties

##### 4.1. X-Ray Spectral Modeling

We perform X-ray spectral modeling using XSPEC (version 12.8.1j; Arnaud 1996) with the C-statistic (cstat) setting, \(^35\) which is more appropriate than \( \chi^2 \) in the low-counts regime (e.g., Nousek & Shue 1989). We group the data (source plus background) from NuSTAR and from other X-ray missions by a minimum of 3 counts and 1 count per bin, respectively, as recommended for use with \( \text{cstat} \). \(^36\)

In all cases, we fit a simple unabsorbed power-law model in order to constrain the effective photon index (\( \Gamma_{\text{eff}} \)) and thus obtain a basic measure of the overall X-ray spectral slope. Figure 3 shows the NuSTAR plus soft X-ray (Chandra, Swift XRT, or XMM-Newton) spectra for the eight extreme NuSTAR serendipitous survey sources, with power-law model fits to each. Flat \( \Gamma_{\text{eff}} \) values (e.g., \( \lesssim 0.5 \)) give empirical evidence for high or CT absorption. Further empirical evidence for CT absorption can be obtained from the detection of a strong fluorescent Fe Kα emission line at \( \approx 6.4 \) keV (with an equivalent width of \( \text{EW}_{\text{Fe Kα}} > 1 \) keV, although lower values do not necessarily rule out CT absorption; e.g., Della Ceca et al. 2008; Gandhi et al. 2017). This reflection feature becomes more prominent with increasing levels of absorption (e.g., Risaliti 2002). To place constraints on \( \text{EW}_{\text{Fe Kα}} \) for our sources, we model the rest-frame \( \approx 4–9 \) keV spectrum as a power law (to fit the continuum) plus an unresolved Gaussian at rest-frame \( E = 6.4 \) keV. For two sources (J0505 and J1512) the emission line is well detected, and \( \text{EW}_{\text{Fe Kα}} \) is therefore constrained. For the remaining six sources the line is undetected, due to insufficient counts, and we report upper limits on \( \text{EW}_{\text{Fe Kα}} \) (assuming a line width of \( \sigma_{\text{line}} = 0.1 \) keV) where the data allow informative constraints. In Table 3 we provide the basic observed X-ray spectral properties for the sample: effective photon indices, Fe Kα line equivalent widths, and observed (i.e., uncorrected for absorption) X-ray luminosities.

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\(^{33}\) http://heasarc.gsfc.nasa.gov/docs/nustar/analysis

\(^{32}\) Froncei et al. (2006); http://cxc.harvard.edu/ciao/index.htm

\(^{33}\) http://www.swift.ac.uk/analysis/xrt/

\(^{34}\) http://xmm.esa.int/sas/

\(^{35}\) The W statistic is actually used, since the background is unmodeled; see http://heasarc.gsfc.nasa.gov/docs/xanadu/xspec/wstat.ps..

\(^{36}\) https://asd.gsfc.nasa.gov/XSPECwiki/low_count_spectra
We use three more spectral models in order to constrain the source properties such as the intrinsic absorbing column density ($N_H$), the intrinsic photon index ($\Gamma$), and the X-ray luminosity. First, we fit a transmission-only model (hereafter the transmission model): a power law attenuated by redshifted photoelectric absorption and Compton scattering of photons out of the line of sight ($\text{CABS} \cdot 2\text{WABS} \cdot \text{PWW}$, in XSPEC formalism). This model represents one extreme of obscured AGN spectra, where the X-ray spectrum is dominated by the primary AGN continuum transmitted directly along the line of sight. Second, we fit a reflection-only model (hereafter the reflection model), which represents a power-law spectrum reflected by circumnuclear material. For this we use the PEXRAV model (Magdziarz & Zdziarski 1995), with the reflection scaling factor set to $-1$ to yield a pure reflection spectrum, and with the other parameters set to default values. This model represents the other extreme of obscured AGN spectra, where the X-ray spectrum is dominated by the reflected AGN continuum, which (in combination with strong Fe line emission) implies very high column densities ($N_H \gg 10^{24}$ cm$^{-2}$). At high column densities, X-ray spectra are typically more complex than the transmission and reflection models above, and ideally any absorbed continuum, reflected continuum, and fluorescent line emission should be modeled in a self-consistent way and assuming a well-motivated geometry. We therefore perform an additional third test using the BNTORUS model (hereafter the torus model; Brightman & Nandra 2011), which was produced using simulations of X-ray radiative transfer through a toroidal distribution of gas. We set the model to an edge-on torus configuration ($\theta_{\text{inclination}}$ and $\theta_{\text{torus}}$ set to $87^\circ$ and $60^\circ$, respectively). In this form, the torus model has the same number of free parameters as the transmission and reflection models and is therefore no less suited to the statistical quality of the data. For every model fit, we account for Galactic absorption with a PHABS multiplicative component, fixed to column density values from Kalberla et al. (2005). In cases where $\Gamma$ and $N_H$ cannot be simultaneously constrained, we fix the intrinsic photon index at $\Gamma = 1.9$ (a typical value for AGNs detected at $3$–$24$ keV; e.g., Alexander et al. 2013; Del Moro et al. 2017). In Table 4 we show the best-fit parameters obtained by applying the three models described above: intrinsic photon indices, column densities, fit statistics, and intrinsic (i.e., absorption-corrected) luminosities.

### Table 2

Summary of the X-Ray Data Adopted for the Spectroscopic and Photometric X-Ray Analyses

| Object (1) | Observation ID (2) | UT Date (3) | $T$ (4) | $S_{\text{net}}$ (5) | $B$ (6) | Observatory (7) | Observation ID (8) | UT Date (9) | $T$ (10) | $S_{\text{net}}$ (11) | $B$ (12) |
|------------|-------------------|-------------|--------|-------------------|--------|----------------|-------------------|-------------|--------|-------------------|--------|
| J0505      | 600610560002      | 2013 Aug 21 | 21.1   | 66                | 53     | XMM-Newton     | 0605090101c      | 2009 Aug 06 | 29.4   | 70                | 46     |
| J0823      | 600610800002      | 2014 Jan 10 | 24.3   | 41                | 67     | XMM-Newton     | 0501210501      | 2007 Oct 14 | 8.4    | 12                | 9      |
| J1410      | 601605710002      | 2015 May 14 | 22.2   | 153               | 125    | Swift XRT      | 00040973002      | 2010 Sep 27 | 5.0    | ...               | ...    |
| J1444      | 901010040002      | 2015 Apr 25 | 38.2   | 62                | 153    | Swift XRT      | 0033768001      | 2015 May 13 | 3.1    | ...               | ...    |
| J1506      | 600612610002      | 2014 Sep 08 | 21.3   | 81                | 122    | Swift XRT      | 00036622001      | 2007 Dec 19 | 9.4    | ...               | ...    |
| J1512      | 60061263002      | 2013 Aug 06 | 13.3   | 153               | 74     | Swift XRT      | 0036622002      | 2007 Dec 21 | 8.7    | ...               | ...    |
| J1534      | 60002026002       | 2013 Aug 13 | 66.7   | 42                | 133    | Chandra        | 00036623001      | 2007 Jun 07 | 6.2    | 11                | ...    |
| J1653      | 60002024002       | 2013 Apr 13 | 18.3   | 14                | 16     | XMM-Newton     | 0652570101c      | 2010 Sep 08 | 43.7   | 73                | 47     |
|            |                   |             |        |                   |        |                | 0652570201c      | 2010 Sep 10 | 44.0   | 82                | 42     |

Notes. Column (1): abbreviated NuSTAR source name. Columns (2) and (3): NuSTAR observation ID and start date (YYYY MM DD). Columns (4), (5), and (6): net exposure time (ks), net source counts, and scaled background counts, respectively, for the extracted 3–24 keV (or 8–24 keV for J1534 and J1653) NuSTAR spectrum. Column (7): soft X-ray observatory with the best (or in some cases, the only) coverage, which we adopt for the analyses. Columns (8) and (9): adopted soft X-ray observation ID(s) and their corresponding start date(s) (YYYY MM DD), respectively. Columns (10), (11), and (12): exposure time (ks), net source counts, and scaled background counts, respectively. For J0505, J0823, J1444, J1512, J1534, and J1653, these columns correspond to the extracted X-ray spectra (at 0.5–10 keV, 0.6–10 keV, and 0.5–8 keV for XMM-Newton, Swift XRT, and Chandra, respectively). For the remaining two sources that are undetected at soft X-ray energies (J1400 and J1506), the Swift XRT data tabulated here are used for photometric constraints.

* Here we use the NuSTAR FPMB data only (i.e., excluding the FPMA data).
* In these cases we limit the NuSTAR spectral analysis to the 8–24 keV band, since the sources are undetected in the soft (3–8 keV) and full (3–24 keV) NuSTAR bands, indicating no significant source emission at <8 keV.
* In these cases we use the combined MOS1+MOS2 data only.
* Here we quote the total exposure time and counts (summing across all observations), since the source is undetected in individual Swift XRT observations.
In one case (J1653) we find that an additional soft-X-ray-dominated model component is necessary to obtain an acceptable fit to the data. For J1653 all three models provide a poor fit to the XMM-Newton plus NuSTAR spectrum (for the transmission, reflection, and torus models, the ratio of the C statistic to the number of degrees of freedom is C/n = 352/200, 311/202, and 335/201, respectively) and leave strong positive residuals at high energies (\( \geq 8 \) keV). This is due to an apparently sudden change in the spectral shape, with the low energies (\( \leq 4 \) keV) dominated by a steep (\( \Gamma \approx 2 \)) component and the higher energies (\( \geq 4 \) keV) dominated by a flatter component (\( \Gamma \approx -0.5 \)). One way to interpret this is an electron-scattered or leaked (due to partial covering) AGN power law at lower energies and a primary AGN continuum penetrating through at higher energies, as is commonly observed for well-studied AGNs in the local universe (e.g., Cappi et al. 2006). The relatively high luminosity (\( L_{0.5-4 \text{keV}} \approx 7 \times 10^{42} \text{ erg s}^{-1} \)) justifies the scattered AGN power-law interpretation rather than, e.g., thermal emission associated with star formation. For J1653 we therefore add an un-obscured power-law component to the three spectral models, with the spectral slope tied to that of the intrinsic AGN power-law continuum. This results in statistically improved fits (see the \( C/n \) values in Table 4) and reasonable scattered power-law fraction constraints (\( f_{\text{scatt}} \approx 0.04\% - 5\% \)).

The source J1534 also shows evidence for a steep soft component in the Chandra spectrum (\( \Gamma_{\text{eff}} \approx 3 \) at 0.5–8 keV), which is dominated by photon counts at \(< 2 \) keV (as described

Figure 3. X-ray spectra in observed count-rate units (top panel for a given source) and in \( E_{\text{F}} \) units (bottom panel for a given source) for the eight extreme NuSTAR sources (Section 4). Black and red correspond to NuSTAR FPMA and FPMB, respectively. The green, blue, and purple spectra represent the available soft X-ray data (as labeled). Letter suffixes (e.g., Swift XRT b) indicate separate observations. See Table 2 for a full description of the adopted data sets. The data are binned to a minimum significance of \( 2 \sigma \) per bin for visual purposes. The \( E_{\text{F}} \) spectra are shown with best-fitting power-law models, binned to match the data (solid lines).
Table 3

Basic X-Ray Spectral Parameters

| Object | $\Gamma_{\text{NuSTAR}}$ | $\Gamma_{\text{Chandra}}$ | EW_{K\alpha} | $L_{2-10}^{\text{obs}}$ | $L_{2-10}^{\text{un}}$ |
|--------|-----------------|-----------------|-----------|-----------------|-----------------|
| J0505  | $-0.1_{-0.8}^{+0.7}$ | $-0.9_{-1.4}^{+1.4}$ | $1.4_{-0.9}^{+1.4}$ | 41.3 | 42.3 |
| J0823  | $0.3_{-1.3}^{+1.3}$ | $1.2_{-0.9}^{+1.2}$ | ... | 42.5 | 44.4 |
| J1410  | $0.3_{-0.4}^{+0.4}$ | ... | ... | 42.0 | 42.7 |
| J1444  | $-0.3_{-0.9}^{+0.9}$ | $0.7_{-1.1}^{+1.1}$ | ... | 44.7 | 45.1 |
| J1506  | $-0.7_{-0.9}^{+0.9}$ | ... | ... | 39.9 | 42.6 |
| J1512  | $0.9_{-0.4}^{+0.4}$ | $-0.6_{-0.6}^{+0.6}$ | $0.76_{-0.26}^{+1.0}$ | 42.4 | 43.2 |
| J1534  | $<0.9$ | $3.3_{-2.4}^{+1.9}$ | ... | 39.8 | 42.7 |
| J1653  | $-0.5_{-0.6}^{+0.9}$ | $2.0_{-0.3}^{+0.3}$ | ... | 42.7 | 44.3 |

Note. Column (1): abbreviated NuSTAR source name. Column (2): NuSTAR effective photon index, i.e., the photon index obtained from approximating the NuSTAR 3–24 keV spectrum as a simple power law. For the sources marked. Column (3): “soft” effective photon index, measured using the available soft X-ray spectra from Chandra, Swift XRT, or XMM-Newton (over the full energy range for the relevant observatory; $\geq 0.5$–10 keV). Column (4): constraint on the Fe K$\alpha$ line equivalent width. Units: keV. Columns (5) and (6): logarithm of the observed (i.e., uncorrected for absorption) X-ray luminosities in the rest-frame 2–10 keV and 10–40 keV bands, respectively. Units: erg s$^{-1}$.

* The constraint was obtained using a combination of NuSTAR and soft X-ray (XMM-Newton or Swift XRT) data, due to weak NuSTAR-only constraints.

in Section 3.1). This is uncharacteristic of pure AGN emission and indicates that at low X-ray energies there is a significant contribution to the spectrum from other radiative processes in the host galaxy. We find that the detection of this soft component is due to the primary AGN spectrum being highly absorbed (see Sections 4.2 and 5) so as not to be well detected by Chandra. Indeed, the AGN is only detectable at $\gtrsim 8$ keV with NuSTAR. The luminosity of the soft X-ray emission ($L_{2-10}^{\text{obs}} = 10^{39.8}$ erg s$^{-1}$; Table 3) is in broad agreement with the expectations for normal galaxy emission based on the X-ray main sequence of star formation (Aird et al. 2017) and given the stellar mass of J1534 ($M_* = 10^{11.1}$ M$_\odot$; based on the spectral energy distribution (SED) modeling in Section 5). If the soft component is instead interpreted as a scattered AGN power law, then the scattered fraction must be small ($f_{\text{scatt}} \lesssim 0.05\%$). For the spectral modeling of J1534 below, we parameterize the steep soft emission with an additional power-law component. We also tested a different approach of simply excluding the $<2$ keV photons, and this yields consistent values for the intrinsic source properties.

For the sources where we model the NuSTAR data simultaneously with soft X-ray (Chandra, Swift XRT, or XMM-Newton) data, there is a general caveat that the soft X-ray observations are not contemporaneous with the NuSTAR data, and AGN variability could thus affect the interpretations. Although highly obscured AGNs such as those presented here show some evidence for lower variability compared to unobscured AGNs (e.g., Awaki et al. 2006), significant variability on year-long timescales is still possible (e.g., Yang et al. 2016; Masini et al. 2017). While our sources generally show no evidence for significant variability (e.g., see the overlapping data in Figure 3), the spectral uncertainties are generally too large to rule out low-level (e.g., factors of $\lesssim 2$) variability. We thus fix the cross-normalization constants to standard values: 1.0 for Chandra:NuSTAR, 1.0 for Swift XRT: NuSTAR, and 0.93 for XMM-Newton:NuSTAR (e.g., Madsen et al. 2015). There is one exception, J0823, where the XMM-Newton:NuSTAR cross-normalization parameter must be left free to obtain statistically acceptable solutions. The transmission and torus models converge to extremely low cross-normalization constants ($\approx 0.01$), and we therefore limit the modeling to the NuSTAR data only. The best-fit reflection model, however, has a less extreme cross-normalization constant of 0.12$^{+0.19}_{-0.08}$ when fitting the XMM-Newton plus NuSTAR data set. The low cross-normalization constants for J0823 may be due to X-ray variability between the 2007 XMM-Newton and the 2014 NuSTAR observations, although we do not draw strong conclusions given the uncertainties for this source.

4.2. Results for the X-Ray Source Properties

Here we summarize the measured X-ray properties. Figure 4 shows the effective photon indices (i.e., the observed spectral slopes) of the sources, as measured with individual X-ray observatories, as a function of X-ray luminosity (uncorrected for absorption). The extreme NuSTAR sources cover a broad range in luminosity. The NuSTAR-measured effective photon indices (right panel of Figure 4) are generally very low (median value of $\Gamma_{\text{eff}} = -0.2$ at 3–24 keV), giving empirical evidence for very high absorption levels. We compare against another sample of extreme systems: highly obscured SDSS-selected Type 2 quasars targeted with NuSTAR (Gandhi et al. 2014; Lansbury et al. 2014, 2015). The two extreme samples cover a similar range of spectral slopes and lie at significantly harder values (i.e., lower $\Gamma_{\text{eff}}$ values) than the general population of “normal” NuSTAR serendipitous survey sources (also shown in Figure 4, for sources with constrained $\Gamma_{\text{eff}}$ values; Lansbury et al. 2017). The measured spectral slopes show a large scatter at soft energies ($\approx 0.5–10$ keV; using Chandra, Swift XRT, and XMM-Newton). For the NuSTAR-observed SDSS Type 2 quasars, this scatter was found to be partly due to an increased contamination at these lower X-ray energies from radiative processes other than the direct AGN emission (e.g., Lansbury et al. 2015), which may also be the case for some of the extreme NuSTAR sources (namely, J1534 and J1653; see Section 4.1). In other words, soft X-ray observations alone would fail to identify 57$^{+11}_{-21}$% of the extreme sources in Figure 4 as highly obscured using spectral slope information (assuming a threshold of $\Gamma_{\text{eff}} = 1$). NuSTAR observations, on the other hand, are highly reliable at identifying the most highly obscured AGNs.

For the purposes of comparing $N_H$ constraints and estimating intrinsic luminosities ($L_X$; shown in Table 4), we adopt the torus model solutions. In one exception (J0823) we adopt the lower-$N_H$ transmission model solution. The adopted best-fitting $N_H$ and $L_X$ values are shown in Figure 5. Based on these intrinsic luminosity constraints, the more distant AGNs ($z > 0.2$) are at “X-ray quasar” luminosities ($L_X \gtrsim 10^{44}$ erg s$^{-1}$), and the less distant AGNs ($z < 0.2$) range from relatively low luminosities up to the quasar threshold ($L_X \approx 10^{42.7}$–$10^{44}$ erg s$^{-1}$). The $N_H$ constraints shown may be conservative for sources where the reflection model gives a statistically acceptable fit to the X-ray spectrum (indicating consistency with $N_H \gtrsim 10^{24}$ cm$^{-2}$). For a similar reason, the Compton-thin constraints shown for J1410 and J1444 may be conservative; the torus modeling also finds statistically acceptable reflection-dominated model solutions at very high CT column densities ($N_H > 6 \times 10^{24}$ cm$^{-2}$) in these cases.
### Table 4
Best-fit Parameters for the X-Ray Spectral Modeling

| Object   | $E$ Range (keV) | $\Gamma_{\text{eff}}$ | C/n | $\Gamma$ | $N_{\text{H}}$ (10$^{24}$ cm$^{-2}$) | C/n | $\Gamma$ | C/n | $N_{\text{H}}$ (10$^{24}$ cm$^{-2}$) | C/n | $L_{2-10}^{\text{in}}$ | $L_{10-40}^{\text{in}}$ | CT |
|----------|-----------------|------------------------|-----|----------|--------------------------------------|-----|----------|-----|--------------------------------------|-----|-------------------------|-------------------------|----|
| J0505    | 0.5–24          | $-0.2 \pm 0.2$         | 164/142 | [1.9]   | $0.87_{-0.37}^{+0.37}$                | 159/139 | 1.3 $\pm 0.4$ | 148/139 | $2.5_{-0.5}^{+0.4}$               | 148/142 | 43.1                     | 42.7                     | Y  |
| J0823    | 0.5–24          | $-0.2 \pm 0.7$         | 78/54 | [1.9]   | $0.73_{-0.21}^{+1.51}$                | 45/33 | 2.6 $\pm 0.5$ | 71/53 | [1.9] | 12.6 $\pm 0.4$ | 41/33 | 44.4                     | 44.4                     | ... |
| J1410    | 3–24            | $0.3 \pm 0.4$          | 78/87 | [1.9]   | $0.74_{-0.03}^{+0.31}$                | 78/87 | 1.8 $\pm 0.4$ | 82/87 | [1.9] | 0.63 $\pm 0.24$ | 80/87 | 43.0                     | 43.0                     | ... |
| J1444    | 0.6–24          | $0.8 \pm 0.5$          | 98/75 | [1.9]   | $0.21_{-0.17}^{+0.28}$                | 104/75 | 2.1 $\pm 0.7$ | 102/75 | [1.9] | 0.24 $\pm 0.28b$ | 103/75 | 45.1                     | 45.1                     | ... |
| J1506    | 3–24            | $-0.7_{-0.16}^{+0.19}$ | 77/64 | [1.9]   | $0.5_{-0.3}^{+0.16}$                 | 82/64 | [1.9] | 79/65 | 1.5 $\pm 0.2$ | 41.2 $\pm 0.33$ | 70/63 | 43.3                     | Y  |
| J1512    | 0.6–24          | $0.4 \pm 0.2$          | 123/98 | [1.9]   | $0.13_{-0.06}^{+0.22}$                | 142/98 | 2.1 $\pm 0.2$ | 112/98 | 2.9 $\pm 0.8$ | 112/97 | 44.0                     | 44.4                     | Y  |
| J1534    | 0.5–24          | $-2.3_{-0.3}^{+0.5}$   | 90/74 | [1.9]   | $2.5_{-0.3}^{+0.2}$                  | 84/72 | [1.9] | 90/73 | 1.6 $\pm 0.4$ | 87/72 | 42.7                     | 42.7                     | y  |
| J1653    | 0.5–24          | $1.9_{-0.3}^{+0.4}$    | 182/194 | 2.3 $\pm 0.5$ | $2.4_{-0.9}^{+0.3}$                  | 165/192 | [1.9] | 179/193 | 2.3 $\pm 0.5$ | 1.6 $\pm 0.15$ | 175/192 | 44.3                     | 44.1                     | y  |

**Notes.** Column (1): abbreviated NuSTAR source name. Column (2): energy range modeled (units of keV). Columns (3) and (4): best-fit results for the unobscured power-law model (pow; also shown in Figure 3), where $\Gamma_{\text{eff}}$ is the power-law photon index. Columns (5)–(12): best-fit results for the transmission, reflection, and torus models, respectively. These include the intrinsic photon index ($\Gamma$; square brackets indicate fixed values), the column density ($N_{\text{H}}$; units of 10$^{24}$ cm$^{-2}$), and the fit statistic ($C/n$, where $C$ is the $C$-statistic and $n$ is the number of degrees of freedom). An error value of +a or −b indicates that the parameter is unconstrained at the upper or lower end. Columns (13) and (14): logarithm of the intrinsic (i.e., absorption-corrected) X-ray luminosities in the rest-frame 2–10 keV and 10–40 keV bands, respectively. Units: erg s$^{-1}$. Column (15): flag to indicate high-confidence CT AGNs and likely CT AGNs (marked as "Y" and "y," respectively). J1653 is marked as "y?" since there is multiwavelength evidence against a CT interpretation (Section 5). For the three sources marked as "-" we cannot strongly rule out CT absorption based on the X-ray modeling.

*a* As detailed in Section 4.1, the transmission and torus model fits for J0823 are performed for the NuSTAR data only (i.e., the XMM-Newton data are excluded).

b For two sources (J1410 and J1444) we show the conservative low-$N_{\text{H}}$ torus model solution in this table, but in each case there is also a second similarly valid solution at very high column densities (for J1410, $N_{\text{H}} > 6 \times 10^{24}$ cm$^{-2}$ and C/n = 92/87; and for J1444, $N_{\text{H}} > 6 \times 10^{24}$ cm$^{-2}$ and C/n = 102/75).

* For J1512, fixing $\Gamma$ to more typical values results in even higher $N_{\text{H}}$ solutions (e.g., a lower limit of $N_{\text{H}} > 8 \times 10^{24}$ cm$^{-2}$ for $\Gamma = 1.9$).
Compton-thin AGN (J1410), one uncertain but likely highly obscured AGN (J0823), and one likely moderately absorbed AGN (J1444). Of the total four likely CT AGNs identified with NuSTAR, none would be identified as CT using just the soft X-ray (<10 keV) data, except possibly J0505, for which the XMM-Newton spectrum alone shows good evidence for a ≳1 keV Fe Kα line.

Prior to this work, only one other AGN has been identified in the NuSTAR extragalactic surveys with strong evidence for CT absorption. This source, ID 330, was identified in the NuSTAR-COSMOS survey (Civano et al. 2015; Zappacosta et al. 2017). Like the robust CT AGNs presented here (J0505, J1506, and J1512), ID 330 lies at low redshift (z = 0.044) and has a high NuSTAR band ratio (see Figure 1). Assuming a BNTORUS-based model to fit the X-ray spectrum, the column density of ID 330 is \( N_H = (1.2 \pm 0.1) \times 10^{24} \text{ cm}^{-2} \) (Civano et al. 2015), which is similar to J0505 and less extreme than J1506 and J1512. Additional CT candidates are identified by Del Moro et al. (2017) and Zappacosta et al. (2017), as part of studies that focus on the broad X-ray spectral properties of NuSTAR extragalactic survey sources. We note that our extreme sample (selected from the total 40-month serendipitous catalog; see Section 2) does not overlap with the Zappacosta et al. (2017) sample, which is a subset of 24 serendipitous sources (plus 39 sources from the NuSTAR dedicated-field surveys).

5. Indirect Absorption Diagnostics

The intrinsic X-ray and MIR luminosities of AGNs are tightly correlated (e.g., Krabbe et al. 2001; Lutz et al. 2004; Horst et al. 2008; Fiore et al. 2009; Gandhi et al. 2009; Lanzuisi et al. 2009; Ichikawa et al. 2012; Matsuta et al. 2012; Asmus et al. 2015; Mateos et al. 2015; Stern 2015; Chen et al. 2017). The observed X-ray-to-MIR luminosity ratio of a source can therefore give an independent, albeit indirect, assessment of the degree of obscuration (e.g., see Alexander 2017, for a recent review); the observed X-ray luminosity for any significantly absorbed AGN will be suppressed with respect to the intrinsic luminosity, causing it to deviate from the X-ray-to-MIR luminosity relation. This diagnostic has been utilized for other NuSTAR studies of obscured AGNs (e.g., Baloković et al. 2014; Lansbury et al. 2014, 2015; Stern et al. 2014; Annuar et al. 2015, 2017; Gandhi et al. 2017; LaMassa et al. 2016).

Figure 6 shows the observed X-ray versus intrinsic 6 \( \mu \)m luminosities for the eight extreme NuSTAR serendipitous survey sources. Adopting the methodology of Assef et al. (2008, 2010, 2013), the AGN \( L_{6\mu m} \) values have been determined using SED modeling of the SDSS and WISE photometry available, where each SED is modeled as the best-fit linear combination of four empirical templates (one AGN template and three different galaxy templates; Assef et al. 2010). The approach allows constraints on the relative contribution of the AGN and the host galaxy to the observed luminosity (see Lansbury et al. 2014, 2015, for applications of the same technique to an SDSS Type 2 quasar sample). For two of the extreme NuSTAR sources (J1444 and J1653) the SED modeling results are consistent with zero contribution from the AGN, and we therefore adopt conservative upper limits for \( L_{6\mu m} \) (Figure 6). For the remaining six sources, the AGN contributes between ≈0.07 and ≈0.77 of the overall luminosity, for the 0.1–30 \( \mu \)m wavelength range (see

Figure 4. Observed X-ray properties: effective photon index (i.e., spectral slope) vs. rest-frame X-ray luminosity (uncorrected for absorption). The left panel shows the properties measured at soft X-ray energies (with Chandra, Swift XRT, or XMM-Newton), and the right panel shows the properties measured at harder X-ray energies with NuSTAR. \( \Gamma_{\text{eff}}^{\text{B}} \) and \( \Gamma_{\text{eff}}^{\text{NuSTAR}} \) are measured for the observed-frame 0.5–10 keV and 3–24 keV bands, respectively. We compare the extreme NuSTAR serendipitous survey sources (black circles, individually labeled) to “normal” serendipitous survey sources (smaller gray circles) and to highly obscured and CT Type 2 quasars which were optically selected and followed up with NuSTAR observations (filled gray squares; Gandhi et al. 2014; Lansbury et al. 2014, 2015).

Figure 5. Rest-frame intrinsic (i.e., absorption-corrected) 10–40 keV X-ray luminosity (\( L_X \)) vs. column density (\( N_H \)), from modeling the X-ray spectra of the extreme NuSTAR serendipitous survey sources (open circles). Each data point corresponds to the torous model solution (except J0823, where the transmission model solution is shown). Following Figure 4, the filled gray squares show a comparison sample of highly obscured Type 2 quasars (Gandhi et al. 2014; Lansbury et al. 2014, 2015). The CT column density region (\( N_H \geq 1.5 \times 10^{23} \text{ cm}^{-2} \)) is highlighted in gray.

Nevertheless, for these two sources we assume the lower-\( N_H \), Compton-thin solutions on the basis that their X-ray-to-MIR luminosity ratios are consistent with those for unobscured AGNs (Section 5).

Considering all of the X-ray spectral constraints together, there are three sources with strong evidence for being CT AGNs (J0505, J1506, and J1512; two of which have supporting evidence from high equivalent width Fe Kα emission, as shown in Table 3), one likely CT AGN (J1534; supporting indirect evidence is presented in Section 5), one possible CT AGN (J1653; although the indirect evidence prefers a lower-obscuration solution; see Section 5), one highly obscured
Figure 6. X-ray luminosities (at rest-frame 2–10 keV and 10–40 keV) vs. rest-frame 6 μm luminosity in νL_ν units (L_6μm). For the data points, we observe X-ray luminosities (i.e., uncorrected for line-of-sight absorption of the X-rays). The extreme NuSTAR serendipitous survey sources are highlighted as orange circles and are individually labeled. We compare to “normal” NuSTAR serendipitous survey sources (smaller blue circles; Lansbury et al. 2017) and to other NuSTAR-observed samples of obscured to CT AGNs (see figure legend). We also compare with known “bona fide” CT AGNs in the local universe (plus signs; distance ≤100 Mpc; data compiled in P. G. Boorman et al. 2017, in preparation), including NGC 1068 and Circinus. The gray regions (with solid borders) highlight the range of luminosity ratios expected in the case of zero X-ray absorption (based on Fiore et al. 2009; Gandhi et al. 2009; Stern 2015; Chen et al. 2017), and the purple regions (with dashed borders) show the approximate X-ray suppression expected for absorption by gas with a column density of N_H = 10^{23} cm^{-2}.

Table 5

SED Modeling Results

| Object | α | L_6μm [10^{42} erg s^{-1}] |
|--------|---|--------------------------|
| J0505  | 0.07 ± 0.05 | 1.5 ± 0.8 |
| J0823  | 0.28 ± 0.08 | 20.3 ± 8.8 |
| J1410  | 0.11 ± 0.07 | 3.0 ± 2.1 |
| J1444  | 0.00±0.19  | <933.2 |
| J1506  | 0.28 ± 0.01 | 11.4 ± 0.7 |
| J1512  | 0.76 ± 0.09 | 36.6 ± 1.7 |
| J1534  | 0.40 ± 0.03 | 35.3 ± 3.8 |
| J1653  | 0.02±0.10  | <26.8 |

Note. Column (1): abbreviated NuSTAR source name. Column (2): fractional contribution of the AGN to the intrinsic luminosity at 0.1 μm–30 μm. Column (3): rest-frame 6 μm luminosity of the AGN.

The resulting uncertainties on L_6μm (also listed in Table 5) are determined from a Monte Carlo resampling of the photometric data over 1000 iterations and are shown in Figure 6.

In Figure 6 we compare with “normal” NuSTAR serendipitous survey sources (Lansbury et al. 2017) and with other NuSTAR-observed highly obscured AGNs, including nearby CT AGNs identified in the NuSTAR snapshot survey (z ≈ 0.01; Baloković et al. 2014), candidate CT Type 2 quasars selected by SDSS (z = 0.05–0.49; Gandhi et al. 2014; Lansbury et al. 2014, 2015), a highly obscured quasar identified in the NuSTAR-ECDFS survey (z ≈ 2; Del Moro et al. 2014), and the CT AGN identified in the NuSTAR-COSMOS survey (z = 0.044; C15). Also plotted are “bona fide” CT AGNs in the local universe (distance ≤100 Mpc; data compiled in P. G. Boorman et al. 2017, in preparation). We compare all sources with the intrinsic X-ray–MIR relation for unobscured AGNs (Fiore et al. 2009; Gandhi et al. 2009; Stern 2015; Chen et al. 2017), and to demonstrate the expected deviation from the relation for highly obscured AGNs, we also show the modified relation for X-ray luminosities suppressed by N_H = 10^{23} cm^{-2} gas. The latter results in a more extreme suppression of the X-ray luminosity for the 2–10 keV band (L_X is decreased by a factor of ≈20) than for the 10–40 keV band (a factor of ≈2 decrease), where the higher-energy photons are less affected by absorption.

For the eight extreme NuSTAR serendipitous survey sources, the X-ray-to-MIR luminosity ratios are in broad agreement with the X-ray spectral modeling results, in that the sources with X-ray spectroscopic evidence for being CT are further offset from the intrinsic L_X–L_MIR relations than the less obscured AGNs. This is especially apparent for J0505, J1506, J1512, and J1534 at 2–10 keV, where these likely CT sources overlap well with the X-ray-to-MIR luminosity ratios of local “bona fide” CT AGNs, as well as luminous highly obscured and CT Type 2 quasars. The L_X–L_MIR ratios are very low in the cases of J1506 and J1534, which appear to lie even lower than local bona fide CT AGNs (including Circinus and NGC 1068), and have observed X-ray luminosities that are suppressed by ≈2–3 orders of magnitude. The X-ray properties of these NuSTAR sources (Section 4.2) suggest that the X-ray weakness is due to extreme absorption, rather than intrinsic X-ray weakness (e.g., Gallagher et al. 2001; Wu et al. 2011; Luo et al. 2014; Teng et al. 2015). J1653 has a relatively high ratio (at both 2–10 keV and 10–40 keV), suggesting a low column density that is in
emission-line properties, we fit the optical spectra for the major lines at rest frame 3500–7000 Å (e.g., [O II], Hβ, [O III], [O I], Hα, [N II], and [S II]) with the pyspeckit software following Berney et al. (2015) and the general procedure in Koss et al. (2017). We correct the narrow-line ratios (Hα/Hβ) assuming an intrinsic ratio of 3.1 and the Cardelli et al. (1989) reddening curve.

For six sources with significantly detected Hα emission lines (signal-to-noise ratio S/N ≥ 4; J0505, J0823, J1410, J1506, J1512, and J1534), the Hα FWHMs range from 269 to 538 km s⁻¹, before correction for instrumental resolution. In no case is a second (broad-line) component required to provide a statistically acceptable fit to the data. These results confirm the visual classifications of these sources as narrow-line systems (Lansbury et al. 2017). We note that J1563 has only a weak detection of Hα, and J1444 is at high redshift (z = 1.539) such that the above emission lines are not in the redshifted spectrum.

For four sources (J0505, J1410, J1506, and J1512), it is possible to apply AGN emission-line diagnostics (e.g., Kewley et al. 2006; Veilleux & Osterbrock 1987) using the [N II]/Hα and [O III]/Hβ emission-line flux ratio constraints. This is not possible for J0823, due to a gap in the spectrum, and for J1534 and J1653, due to the low S/N of the key emission lines. Figure 7 shows the location of the former four sources on the Baldwin–Phillips–Terlevich (BPT) diagram. All four sources fall into the AGN region based on the upper limits for the Hβ line, which is weak to undetected (S/N < 3). The weak Hβ line emission is likely due to extinction by dusty gas and has previously been observed for X-ray-selected obscured AGNs, particularly in mergers (e.g., Koss et al. 2016a, 2016b). We also note that Hβ is undetected for J0823, J1534, and J1653, and even [O III] is undetected in the case of J1534. The seven z < 0.4 extreme NuSTAR AGNs would thus be unidentified in any optical surveys requiring the detection of Hβ.

6.2. Host Galaxies

The five lower-redshift (z < 0.2) extreme NuSTAR sources (J0505, J1410, J1506, J1512, and J1534) have well-resolved host galaxies at optical wavelengths, while the higher-redshift sources are consistent with point-source emission. Four of the five lower-redshift sources are likely CT systems based on our X-ray analyses and also have relatively high quality optical coverage from Pan-STARRS, which appears consistent with being a companion galaxy to 2MFGC 04170. The other lower-redshift source (J1410), on the other hand, is Compton-thin and is limited to low-quality optical coverage from photographic plate observations. Here we comment on the host galaxies, and nearby companion galaxies, for the lower-redshift sources.

J0505.—The optical counterpart is 2MFGC 04170, a highly inclined disk galaxy. The Pan-STARRS coverage of 2MFGC 04170 reveals spatially extended emission at ≈12° offset (or a projected separation of ≈9 kpc) and at a position angle of ≈70°, which appears consistent with being a companion galaxy to 2MFGC 04170 (see Figure 8). We refer hereafter to this second companion source as J050601.2–235002.6. Since this source had no available redshift information, we performed follow-up spectroscopy with Keck (provided in the Appendix). We find that J050601.2–235002.6 lies at z = 0.137 and is therefore a...
background galaxy that is coincidentally aligned along the line of sight, rather than being a merging companion to 2MFGC 04170.

J1506.—The optical counterpart is UGC 09710, an edge-on Sb spiral galaxy belonging to a close spiral–spiral galaxy pair in an early-stage major merger (see Figure 8), and separated from its similar mass partner galaxy (IC 1087; \( z = 0.035; S0-a \) type) by \( \approx 16 \) kpc in projection (Yuan et al. 2012). Physical disturbances resulting from the major merger could potentially be related to an increase in the central gas content. In the Appendix we present a Palomar optical spectrum for the companion galaxy (IC 1087), which shows a possible AGN (also consistent with a LINER classification) with a dominant galaxy continuum. \([\text{O} \text{III}]\) and \( \text{H}\beta \) are undetected for the companion galaxy (presumably due to host galaxy dilution), and the \([\text{N} \text{II}]\):\( \text{H}\alpha \) line strength ratio is very high, but is likely affected by stellar absorption. For this companion galaxy, there is no additional evidence from the \textit{WISE} colors for an AGN, and the source is undetected in the current X-ray coverage.

J1410.—The available photographic plate coverage (from the UK Schmidt Telescope) shows an extended host galaxy, but the low data quality precludes type and disturbance classifications. Nevertheless, there do not appear to be any nearby (massive) companion galaxies.

J1512.—We have obtained \( R \)-band imaging with the ESO-NTT (shown in Figure 8), which is in visual agreement with the host being a relatively undisturbed early-type galaxy. The neighboring optical sources are consistent with being unresolved point sources, with FWHMs similar to the seeing (\( \approx 1''5 \)), and are therefore unlikely to be associated with J1512.

J1534.—The Pan-STARRS imaging (Figure 8) shows good evidence that the optical host galaxy (SDSS J153445.80 +233121.2; \( z = 0.160 \)) is undergoing a major merger with a narrowly offset companion galaxy (SDSS J153446.19 +233127.1; no spec-z); the respective galaxy nuclei are separated by \( \approx 8'' \) (or \( \approx 22 \) kpc in projection), and likely extended tidal features are visible. The merger stage is not clear. We present Palomar spectroscopic follow-up for the companion galaxy in the Appendix, although there are no significantly detected emission or absorption features.

A notable feature of the galaxies is that both J0505 and J1506 have close to edge-on geometries, which could contribute at least some of the observed X-ray obscuration. The axis ratios of the host galaxies are \( b/a = 0.24 \) and 0.23 for J0505 and J1506, respectively, based on isophotal fitting of the galaxy images in Figure 8 (using the IRAF task \texttt{ellipse}). The remaining two likely CT sources (J1512 and J1534), on the other hand, have axis ratios exceeding \( b/a = 0.6 \). Although the source numbers are currently small, the above implies a relatively high fraction (50% ± 33%) of close to edge-on systems for CT AGNs selected by \textit{NuSTAR}. For comparison, only \( \approx 16\% \) of the general hard-X-ray-selected AGN population have \( b/a < 0.3 \), based on isophotal analyses for the \textit{Swift} BAT AGN sample (Koss et al. 2011). Although the difference is only weakly significant, a similar result has also been reported for CT AGNs selected with \textit{Swift} BAT (Koss et al. 2016a). Other studies, however, find that edge-on galaxy inclinations are not clearly related to CT absorption (e.g., Anuar et al. 2017; Buchner & Bauer 2017).
6.2.1. A High Fraction of Galaxy Mergers for the Compton-thick AGNs?

It is interesting that two of the four likely CT AGNs (J0505, J1506, J1512, and J1534) are hosted by galaxy major mergers (see Figure 8). To assess the statistical significance of the apparently high merger fraction for these extreme NuSTAR serendipitous survey AGNs ($f_{\text{merger}} = 50 \pm 33\%$; the errors represent binomial uncertainties), we can search for similar merging systems in the sample of nonextreme (or “normal”) serendipitous survey AGNs. To this end, from the overall serendipitous survey sample, we apply a cut of $BR_{\text{Nu}} < 1.7$, thus limiting to those sources that do not have very hard NuSTAR spectra (based on the $BR_{\text{Nu}}$ threshold in Section 2). We limit this comparison sample to source redshifts of $0.01 < z < 0.2$, thus matching the redshift range of the four extreme sources. We exclude two sources from the sample that are likely strongly associated with the science targets of their NuSTAR observations (similar to the exclusion of J2028 from the extreme sample; see Section 2). These cuts leave 36 normal NuSTAR sources. Finally, we limit the sample to the 26 (out of 36) sources that are covered by Pan-STARRS observations and therefore have optical coverage that is of comparable quality to the four extreme NuSTAR sources. As a result, the comparison of visual merger classifications between the two different samples is unlikely to be significantly affected by variations in optical imaging sensitivity. The comparison sample is matched in X-ray luminosity distribution to the extreme NuSTAR AGNs (with a Kolmogorov–Smirnov test p-value of 0.8).

Of the 26 normal AGNs, we identify one that has evidence for a galaxy major merger, with a comparably sized companion galaxy lying at the same redshift and offset by a projected distance of $\approx 25$ kpc. There are an additional two normal AGNs with possible evidence for mergers, although the candidate companion galaxies are relatively small in size, with unknown redshifts. We conservatively assume that two of the normal AGNs are in major mergers with $<30$ kpc separation companions.

Our estimate for the major-merger fraction of normal NuSTAR AGNs is therefore $f_{\text{merger}} = 8^{+12}_{-5}\%$. This is in agreement with the ($<30$ kpc separation) major-merger fraction for Swift BAT AGNs ($f_{\text{merger}} = 13^{+2}_{-3}\%$; Koss et al. 2010). Figure 9 compares the above merger fractions. We additionally compare with low-redshift inactive galaxies and optical Type 2 AGNs (both from the SDSS), which are matched to the Swift BAT sample (Koss et al. 2010) and have very low merger fractions compared to the Swift BAT and extreme NuSTAR AGNs. At low significance levels of $1.8\sigma$ and $1.7\sigma$ (according to the Fisher exact probability test), the extreme (very hard, CT) NuSTAR AGNs have a higher merger fraction than both the normal NuSTAR AGNs and the Swift BAT AGNs, respectively. This could be a result of Compton-thick phases of black hole growth being more strongly linked (than less obscured phases) to the merger stage of the galaxy evolutionary sequence.

The above result is of interest given recent findings for other AGN samples. Kocevski et al. (2015) find evidence that highly obscured ($N_{\text{H}} \gtrsim 3 \times 10^{23}$ cm$^{-2}$) AGNs at $z \sim 1$ have a higher frequency of merger/interaction morphologies relative to less obscured AGNs matched in redshift and luminosity. Furthermore, Koss et al. (2016a) noted a high close ($<10$ kpc) merger fraction for likely CT Swift BAT AGNs at $z \lesssim 0.03$ ($f_{\text{merger}} = 22\%$; i.e., 2/9). The recent study of Ricci et al. (2017) indicates a possible connection between the late stages of galaxy mergers and high AGN obscuration, in a sample of local luminous and ultraluminous infrared galaxies (U/LIRGS), using a combination of dedicated and archival X-ray observations. Taken together, the results may suggest a departure from simple orientation-based unified models of AGN obscuration and indicate an evolutionary scenario where highly obscured phases of black hole growth can be associated with a merger-driven increase in the circumnuclear gas content (e.g., Sanders et al. 1988; Draper & Ballantyne 2010; Treister et al. 2010). An increased sample size and deeper imaging would help to further test the CT AGN–merger connection using the NuSTAR serendipitous survey.

Figure 10. Top panel: observed cumulative number counts (and 90% CL uncertainties), as a function of $8-24$ keV flux ($S_{\text{ex}}$), for the CT AGNs identified in the NuSTAR serendipitous survey. The gray circles show the number counts for all four CT AGNs. The black square shows the modified number counts when removing the three low-redshift CT AGNs (J0505, J1506, and J1512; see Section 7). We compare to predicted tracks for CT AGNs (dashed lines) and all AGNs (solid lines) based on the models of A15, U14, B11, and T09. The dotted lines show modifications of the CT model tracks to account for the spectroscopic incompleteness of the serendipitous survey. Bottom panel: “intrinsic” cumulative number density (and 68% CL uncertainties) as a function of flux.

7. The Prevalence of Compton-thick Absorption

We have taken advantage of the relatively large sample size of the NuSTAR serendipitous survey to identify rare highly obscured AGNs. While all of the eight extreme sources investigated are consistent with being highly obscured, four in
particular are likely CT (J0505, J1506, J1512, and J1534). A fifth source (J1653) is a CT candidate based on the X-ray analysis, but this result is in tension with the indirect constraints (see Section 5). Here we assess how the observed number of CT AGNs in the NuSTAR serendipitous survey compares with the number expected from AGN population models, which are informed by the results from previous (primarily <10 keV) X-ray surveys. We consider the hard-band (8–24 keV) selected serendipitous survey sample, since this is the energy band in which NuSTAR is uniquely sensitive, and Galactic latitudes of $|b| > 10^\circ$ (i.e., out of the Galactic plane). We conservatively exclude J1653. The top panel of Figure 10 shows the observed (cumulative) number of CT sources as a function of limiting flux, and these results are compared to model predictions for the observed numbers of CT AGNs and all AGNs. For these predictions, we fold the area-sensitivity curve of the serendipitous survey through models for the evolution of the X-ray luminosity function (XLF) and the $N_\text{H}$ distribution of AGNs, from Treister et al. (2009, hereafter T09), Ueda et al. (2014, hereafter U14), Aird et al. (2015, hereafter A15), and the updated version of Ballantyne et al. (2011, hereafter B11). The updates to the B11 model are summarized in Harrison et al. (2016). We additionally show, in the bottom panel of Figure 10, the “intrinsic” cumulative number densities (i.e., the sky number counts before accounting for the survey sensitivity; $N(>S)$, in units of deg$^{-2}$).

In Figure 10 the gray circle data points show the number counts for all four CT AGNs. There is an apparent excess in the CT number counts at high fluxes, compared to the model predictions. This excess may be expected given that the three lowest-redshift, highest-flux sources (J0505, J1506, and J1512; $z < 0.07$) show evidence for being weakly associated with the Swift BAT AGN targets of the NuSTAR observations (see Section 2.1), and also given that galaxy clustering tends to be high around BAT AGNs (e.g., Cappelluti et al. 2010; Koss et al. 2010). In Figure 10 we also show the CT number counts using J1534 only (i.e., excluding J0505, J1506, and J1512; black square data point). Although not particularly constraining, this brings the number counts into better agreement with all of the models (T09, B11, U14, A15, and Gilli et al. 2007), suggesting consistency with a wide range of intrinsic CT fractions ranging from $f_{\text{CT}} \approx 10\%$–40\%, at least for $z > 0.07$. For comparison, Zappacosta et al. (2017) study the X-ray spectral properties of NuSTAR extragalactic survey sources and find that the range of CT fractions allowed by their sample is broad ($f_{\text{CT}} \approx 10\%$–70\%). The NuSTAR survey constraints on $f_{\text{CT}}$ are therefore in broad agreement with $z \geq 0.1$ constraints from soft (<10 keV) X-ray observatories ($f_{\text{CT}} \approx 30\%$–50\%; e.g., Brightman & Ueda 2012; Brightman et al. 2014; Buchner et al. 2015).

However, it is important to consider independently the low-redshift ($z < 0.07$) regime, where we have detected the highest numbers of CT AGNs. Although the overall number counts in this regime may have an upward excess with respect to model predictions (as mentioned above), the CT fraction should be unaffected. The observed CT fraction for the $z < 0.07$ NuSTAR serendipitous survey sample is $f_{\text{CT}}^\text{obs} = 50\pm19\%$ (68\% CL binomial uncertainties). The intrinsic X-ray luminosity range of this subsample is $41.3 < \log(L_{0.1-40\text{keV}}) < 44.0$. Figure 11 compares our observational constraint to model predictions as a function of 8–24 keV flux. We find a higher CT fraction than is expected from the models. The difference is statistically significant in one case (>3$\sigma$; compared to A15) and at lower significance levels for the remaining models (<3$\sigma$; comparing to T09, B11, and U14). In Figure 11 we additionally compare with data points for the higher-flux Swift BAT survey (Burlon et al. 2011; Ricci et al. 2015), for which we have converted to the 8–24 keV NuSTAR band assuming $\Gamma_{\text{IB}} = 1.9$. At present, the origin of the high observed CT fraction at $z < 0.07$ is unclear. A likely explanation is that the current models are not well constrained for the new parameter space probed with NuSTAR, in which case the AGN population models require updating. An alternative possibility, however, is that $f_{\text{CT}}^\text{obs}$ is boosted owing to a real connection between CT absorption and the large-scale environment, in combination with NuSTAR having preferentially targeted (at $z < 0.07$) fields with relatively high galaxy densities (e.g., fields around Swift BAT AGNs).

Finally, we note that the number of CT AGNs presented here could be a lower limit to the total number within the NuSTAR serendipitous survey as there are additional sources, not included in this work, that have band-ratio limits consistent with a large range in column density (e.g., see Figure 1), and any CT sources with relatively soft spectral shapes could potentially be missed by our initial selection (Section 2). Alternative approaches (e.g., detailed X-ray or multiwavelength analyses of the broader sample) may tease out additional CT AGNs within the sample. However, large improvements on the constraints presented here will require further survey data from sensitive hard X-ray missions. Further data will be provided by the continued NuSTAR operations, which are likely to increase the serendipitous sample to $\geq$1000 sources, and potentially by future high-sensitivity >10 keV observatories (e.g., the High-Energy X-ray Probe, or HEX-P, mission concept currently under study; PI F. Harrison; see Brandt & Alexander 2015, for a brief overview).

58 The CT fraction is defined here as the fraction of all AGNs that are CT.
8. Summary

In this paper we have searched for the most extreme sources in the NuSTAR serendipitous survey, in terms of having very hard spectral slopes (BR_\text{Nu} \geq 1.7). The eight selected sources are all candidates for being highly obscured AGNs. A detailed look at the broadband (0.5–24 keV) X-ray data available, as well as the multiwavelength properties of these sources, has yielded the following main results.

1. The X-ray spectral analyses find that three of the extreme NuSTAR sources (J0505, J1506, and J1512) are newly identified robust Compton-thick (CT) AGNs at low redshift (z < 0.1). An additional source at higher redshift (J1534) is likely CT. The remaining four extreme sources are consistent with being CT or at least moderately absorbed; see Section 4.2.

2. Most (three out of four) of the likely CT AGNs identified with NuSTAR would not have been identified as highly obscured systems based on the low-energy (<10 keV) X-ray coverage alone. J1506 is a notable example: a newly uncovered CT AGN in the nearby universe (z = 0.034; \text{N}_\text{H} > 2 \times 10^{24} \text{ cm}^{-2}; L_X \approx 2 \times 10^{43} \text{ erg s}^{-1}), hosted by a previously known galaxy major merger; see Sections 4.2 and 6.2.

3. For all eight extreme sources, the optical spectra show evidence for narrow-line AGNs or galaxy-dominated spectra, supporting the X-ray classifications as obscured and CT AGNs; see Section 6.1. Measurements of the X-ray-to-MIR luminosity ratio, an indirect absorption diagnostic, are also broadly congruent with the X-ray classifications. Two sources (J1506 and J1534) have particularly extreme ratios, lying even lower in L_X/L\text{MIR} than the well-known CT AGNs in the local universe; see Section 5.

4. A high fraction (50% ± 33%) of the likely CT AGNs are hosted by galaxy major mergers. This is higher than the major-merger fractions for “normal” NuSTAR serendipitous survey sources and for Swift BAT AGNs, at a low significance level, motivating larger future studies; see Section 6.2.

5. We estimate the number counts of CT AGNs for the hardband (8–24 keV) selected serendipitous survey sample at |b| > 10°. The number counts are broadly harmonious with AGN population models over the main redshift range of the survey (0.1 \approx z \approx 2), but there is disagreement at low redshifts (z < 0.07) where we find evidence for a high observed CT fraction of n_{\text{obs}} = 30\% ± 12\%; see Section 7.

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Facilities: Chandra, ESO La Silla, Keck, Magellan, NuSTAR, Palomar, Pan-STARRS, SDSS, Swift, WISE, XMM-Newton.

Appendix

A.1. Optical Spectra for the Extremely Hard NuSTAR Serendipitous Survey Sources

Here we provide the optical spectra (Figure 12) for the eight extreme NuSTAR AGNs, which are discussed in Section 6.1. The identified emission and absorption lines are highlighted in Figure 12 and are tabulated in Appendix A.2 of Lansbury et al. (2017).

A.2. Optical Spectra for Companion Galaxies

A.2.1. J0505

As described in the main text, with the Keck telescope we performed optical spectroscopy for J050601.2–235002.6, the apparent companion galaxy to 2MFGC 04170 (the host galaxy for J0505). The resulting spectrum is shown in Figure 13. The relatively high redshift (z = 0.137) confirms that this is a background galaxy and a chance alignment with 2MFGC 04170 (z = 0.036).

A.2.2. J1506

As described in the main text, J1506 belongs to one of two galaxies in a major merger. With the Hale telescope at Palomar Observatory we performed optical spectroscopy for the companion galaxy (known as IC 1087). The resulting spectrum is shown in Figure 14.

A.2.3. J1534

As described in the main text, J1534 (hosted by galaxy SDSS J153445.80+233121.2) appears to be undergoing a major merger with a neighboring galaxy (SDSS J153446.19+233127.1). Since no spectroscopic redshift is available for the latter galaxy, we performed optical spectroscopy with the Hale telescope at Palomar Observatory, the spectrum from which is shown in Figure 15. Since no clear emission or absorption features are detected, this companion requires deeper spectroscopic observations in the future to reliably determine the redshift.
Figure 12. Optical spectra for the extreme NuSTAR serendipitous survey sources. The horizontal axis shows the observed-frame wavelength in units of Å. The vertical axis shows the flux ($f$) in units of erg s$^{-1}$ cm$^{-2}$ Hz$^{-1}$ for all sources except J0505 and J1410, for which the vertical axis shows the counts. The vertical dashed gray lines mark the emission and absorption lines identified.
Figure 13. Keck optical spectrum for J05061.2–235002.6, the apparent companion galaxy to 2MPGC 04170 (the host galaxy for J0505). Multiple emission and absorption lines are identified and labeled here.

Figure 14. Palomar optical spectrum for IC 1087, the merging companion galaxy to 2MPGC 04170 (the host galaxy for our lowest-redshift extreme NuSTAR source, J1506). Multiple emission and absorption lines are identified and labeled here.

Figure 15. Palomar optical spectrum for SDSS J153446.19+233127.1, the merging companion galaxy to SDSS J153445.80+233121.2 (the host galaxy for J1534). The continuum is detected, although no clear emission or absorption lines are identified, precluding a spectroscopic redshift measurement.

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