Discovery of two pulsars towards the Galactic Centre

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ABSTRACT
We report the discovery of two highly dispersed pulsars in the direction of the Galactic Centre made during a survey at 3.1 GHz with the Parkes radio telescope. Both PSRs J1745−2912 and J1746−2856 have an angular separation from the Galactic Centre of less than 0.3° and dispersion measures in excess of 1100 cm\(^{-3}\)pc, placing them in the top 10 pulsars when ranked on this value. The frequency dependence of the scatter-broadening in PSR J1746−2856 is much shallower than expected from simple theory. We believe it likely that the pulsars are located between 150 and 500 pc from the Galactic Centre on the near side, and are part of an excess population of neutron stars associated with the Centre itself. A second survey made at 8.4 GHz did not detect any pulsars. This implies either that there are not many bright, long-period pulsars at the Galactic Centre or that the scattering is more severe at high frequencies than current models would suggest.

Key words: pulsars:general, pulsars:individual:J1745−2912;J1746−2856

1 INTRODUCTION

Surveys for radio pulsars have been extremely successful over the past decade (Lorimer 2005). All-sky surveys and deep surveys of the Galactic plane have doubled the total number of known radio pulsars to above 1700. At the same time, targeted surveys have discovered very young pulsars in supernova remnants, old millisecond pulsars in globular clusters and a population in the Small and Large Magellanic Clouds. In spite of these successes, no pulsars are known within ∼1° of the Galactic Centre (GC).

Following the Johnston et al. (1995) targeted search of the GC at 1500 MHz (no detections), the most sensitive search at 1500 MHz is the Parkes multibeam survey which integrated for 35 mins over a large fraction of the Galactic plane including the GC. That survey discovered PSR J1747−2802, which has a period of 2.8 s, the smallest angular separation from the GC of all the known pulsars (∼1°) and a high dispersion measure (DM) of 835 cm\(^{-3}\)pc (Morris et al. 2002). At higher frequencies, the most sensitive survey was undertaken with the Effelsberg telescope which surveyed a small region (0.2 degrees or 32 pc radius) around the GC at 4850 MHz (Kramer et al. 2000; Klein 2004). In spite of a sensitivity below 100 µJy and the high observing frequency the survey failed to discover any new pulsars.

Sgr A* is at the heart of our Galaxy and is believed to be a black hole with a mass of \(3 \times 10^6 M_\odot\) (Eckart et al. 2002) at a distance of ∼8 kpc. A nearby stellar cluster appears to contain early-type stars with masses of \(10 – 20 M_\odot\). Two of the stars in this cluster are in highly eccentric orbits about the black hole with periods of 15 and 30 years. It is likely that shorter orbits exist, however confusion in the infra-red has so far prevented their detection. The presence of these high mass stars and young supernova remnants is a good indicator that active radio pulsars are also likely to exist close to the GC. Pfahl & Loeb (2004) estimate that there are 1000 radio pulsars within a few pc of the GC, a small fraction of which are potentially detectable using current instruments provided that scatter broadening does not render them invisible at any sensible observing frequency (see also Cordes & Lazio 1997). However, it is likely that the entire GC volume out to ∼200 pc contains an overabundance of pulsars generally. Estimates show that about 10% of all the high mass stars (the pulsar progenitors) in the Galaxy are contained within this volume (Figer et al. 2004).

One can therefore expect to detect a population of pulsars which would be extremely useful in probing the GC and its conditions: their number and age distribution would probe the past star formation history (Hartmann 1995); their period derivatives can constrain the gravitational potential in the GC; pulsar timing would enable us to probe the space-time around the super-massive black hole in the GC due to a variety of relativistic effects (Wex & Kopeikin...
2 OBSERVATIONS AND DATA REDUCTION

The surveys were carried out using the Parkes radio telescope at centre frequencies of 3.1 and 8.4 GHz. At the lower frequency the total bandwidth was 576 MHz in each of two polarizations which was subdivided into 192 channels each of width 3 MHz. At 8.4 GHz, 288 channels were employed for a total bandwidth of 864 MHz. The outputs from the channels were 1 bit digitised and sampled every 250 μs and subsequently written to disk for off-line processing. Figure 1 shows the surveys areas superposed on a continuum image of the GC region.

Observations at 3.1 GHz were carried out from 2005 July 19 to 22. A total of 32 pointings were made, each with an integration time of 70 mins. At this frequency, the half-power width of the telescope beam is 7 arcmin. The survey therefore covered 0.34 square degrees on the sky or a box approximately 90 pc across at the distance of the GC. On cold sky, the system equivalent flux density was ~45 Jy, as measured through observations of the calibrator source Hydra A. However, conditions near the GC contribute substantially to this value. From the maps of Reich et al. (1984), we estimate a contribution of ~25 Jy in the outer regions of the survey, ~65 Jy in the inner regions and up to ~550 Jy at the GC itself. For pulsars with a duty cycle of 10%, the detection threshold (10-σ) is then ~120 μJy (outer regions), ~190 μJy (inner regions) and ~1 mJy at the GC.

Observations at 8.4 GHz were carried out from 2005 September 13 to 16. The survey involved 31 pointings, each observed for 70 mins. The half-power width of the telescope beam is 2.4 arcmin and the survey covered 0.04 square degrees on the sky. The system equivalent flux density on cold sky was 48 Jy. Additions to this from emision at the GC were estimated from the maps of Seiradakis et al. (1989) to be ~4 Jy in the outer regions of the survey, ~10 Jy in the inner regions and ~100 Jy at the GC. The 10-σ detection threshold is then ~70 μJy (outer regions), ~80 μJy (inner regions) and ~200 μJy at the GC.

Data reduction was carried out using the sigproc* software package. An initial pass through the data involved resampling to 1 ms and applying 415 (at 3.1 GHz) or 62 (at 8.4 GHz) trial dispersion delays to the data for a range of dispersion measures (DMs) up to 10000 cm−3pc. A 22− point fast fourier transform was then done on the dedispersed time series and the resultant power spectrum searched for significant spikes. Harmonic summing in 4 stages up to a factor of 16 was carried out and the most significant signals written to disk. These were then time-folded to produce a candidate pulse profiles for subsequent visual inspection.

A search for isolated dispersed bursts of emission with signal-to-noise ratios above a 5-σ threshold was also performed (McLaughlin et al. 2006). Time series were smoothed with boxcar of various widths to increase our sensitivity to broadened pulses. No sources of bursts were found, though the high frequency of these surveys and hence the relatively low dispersion delay, makes distinguishing radio frequency interference from signals of astrophysical origin difficult.

3 TWO NEW PULSARS

Two periodicities, near 945 and 187 ms, were stand-out candidates from the data reduction of the 3.1 GHz data, with signal to noise ratios of 41.5 and 9.4. Confirmation of the first pulsar, PSR J1746–2856, came from analysis of archival data from the Parkes multi-beam survey at 1.4 GHz. The 945 ms pulsar is highly scattered at that frequency but is clearly detected with a signal to noise ratio of 16. A re-detection at 3.1 GHz was made at Parkes on 2005 August 26. The second pulsar, PSR J1745–2912, was not seen in the archival 1.4 GHz data but was confirmed at Parkes at 3.1 GHz on 2005 August 27. Successful detection of both pulsars at 5 GHz was also made using the Effelsberg telescope in early 2005 September, and more accurate positions were obtained by performing a grid search around the discovery locations. The location and pulse profiles for both pulsars are shown in Figure 1.

The pulsars have very large DMs; only 14 pulsars were previously known with DMs in excess of 1000 cm−3pc and of these only PSR B1758–23 is within 15 degrees of the GC. Scattering times were estimated using a technique described in Lohmer et al. (2001). We assumed that there is

| Table 1. Parameters for PSRs J1745−2912 and J1746−2856 |
|-------------------------------|-------------------|
| R.A. (J2000)                  | Epoch             |
| 17h 45m 50(10)               | 53609.30          |
| Dec (J2000)                  |                  |
| −29° 12′ (2)                 | 53609.30          |
| Gal. longitude (deg)         | Epoch             |
| 359.79                       | 53704.50         |
| Gal. latitude (deg)          | Epoch             |
| −0.18                        | 53704.50         |
| P (ms)                       | Epoch             |
| 187.3794(2)                  | 53704.50         |
| DM (cm−3pc)                  | Epoch             |
| 1130(20)                     | 53704.50         |
| S1.1 (mJy)                   | Epoch             |
| 0.12                         | 53704.50         |
| Age (Myr)                    | Epoch             |
| 1.2                          | 53704.50         |
| B (×1012 G)                  | Epoch             |
| 3.5                          | 53704.50         |

* http://sigproc.sourceforge.net
The scattering time is $15 \pm 2$ ms at 1.4 GHz, and the measured values of the DM and the scattering time are $\sim 50$ ms (at 1 GHz). Following the confirmation of the pulsars, subsequent timing observations of PSR J1746–2856 were carried out using the Lovell telescope at the Jodrell Bank Observatory at a frequency of 1.4 GHz with additional data at 3.1 GHz from the Parkes telescope. A total of 49 observations since mid-2005 have allowed us to determine a timing solution for this pulsar. As yet, a timing solution has not been obtained for PSR J1745–2912 due mainly to its lack of detectability at 1.4 GHz. Parameters for both pulsars are listed in Table 1.

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The two pulsars are both at an angular separation of $0.3^\circ$ from the GC, corresponding to $\sim 40$ pc lateral displacement. If they had a similar radial displacement from the GC, they would be well inside the putative scattering screen. Their measured scatter broadening values of $\sim 500$ ms would therefore be about two orders of magnitude lower than expected. This discrepancy is made even worse if the pulsars are located behind the GC. Furthermore, the measured DMs are a factor of 2 less than expected for the GC. We can also examine the scatter broadening of OH maser spots in the GC. The two OH masers nearest PSR J1745–2912 are OH359.762+0.120 and OH359.880–0.087 with angular sizes of 1700 and 2800 mas (scaled to 1 GHz) respectively (van Langevelde et al. 1992; Frail et al. 1994). We can convert our measured time scatter broadening values of $\sim 750$ ms at 1 GHz to an angular broadening and obtain a value of 700 mas assuming a distance of $8.5$ kpc and a scattering screen located close to the pulsar. Again this likely indicates the pulsars are in front of the GC.

Both the Taylor & Cordes (1993) and Cordes & Lazio (2002) models of the electron density consist of an outer, thick disk component, an inner, annular component, and spiral arms. Although the models differ significantly in some parts of the Galaxy, they give very similar results in the inner $\sim 1$ kpc, largely due to the lack of constraints there. If the pulsars are indeed in front of the scattering screen but within the inner few hundred pc, the expected DM would then be only $\sim 650$ cm$^{-3}$ pc in both of the models listed above. Scattering is expected to be $\sim 50$ ms at 1 GHz. The measured values of the DM and the scattering time are significantly higher than the models would suggest, perhaps indicating that a filled centre model, rather than an annular one may be more appropriate.

In any case, it seems most likely that the pulsars would have to be located in front of the scattering screen and at least 150 pc from the GC. This still leaves the question as to their place of birth; in the sections below we consider possible origins for the birth location of these pulsars.
5 LOCATION AND ORIGIN OF THE NEWLY DISCOVERED PULSARS

5.1 Field pulsars: Formation at $D > 1$ kpc from the GC

There is a possibility that these pulsars are merely field pulsars on the near side of the GC but not associated with any (excess) neutron star population there. However, the 3.1 GHz survey would not expect to find any pulsars from the normal field population for two reasons. First, the survey area is tiny; the Parkes multibeam survey detected pulsars at a rate of $\sim 1$ per square degree in the inner Galaxy whereas our survey covered only 0.34 square degrees with a roughly similar effective sensitivity (taking into account the survey parameters, the background temperature and the spectral index of pulsars). Secondly, there is good evidence that the number of field pulsars per unit area actually decreases close to the GC (Johnston 1994; Yusifov & Küçük 2004; Lorimer 2004) making detections even more unlikely in this volume.

We have performed a simulation of the normal Galactic population of pulsars using the $z$-height and radial distributions of Johnston (1994) and the space velocity distribution of Lyne & Lorimer (1994) with a random orientation of the birth kick. We allow a random (flat) distribution of ages up to a maximum of $10^7$ yr and let the pulsars evolve in the Galactic potential. In this simulation we find only 0.1% of the pulsars are located in the region covered by the 3.1 GHz survey. In contrast, if we simulate a population of GC neutron stars with a Gaussian radial distribution with $\sigma = 70$ pc, we find that some 10% of the sample remains inside the survey region. There are estimates that about 10% of all high mass star formation takes place in the inner few hundred pc of the Galaxy. If this is the case then, in the region covered by the 3.1 GHz survey, one would expect to have 10 times more pulsars with origins from the GC population compared to those pulsars which have migrated inwards from the outer Galaxy.

The similarities in the parameters of these two pulsars are striking. Their DMs and scattering times are similar, and they have a similar angular separation from the GC. This seems unlikely to be random chance. Taking all this information together, it therefore seems highly probable that both the newly discovered pulsars are part of an ‘extra’ population directly linked to conditions at the GC.

5.2 Formation within $500$ pc of the GC

The pulsars could have been born in the stellar cluster which surrounds Sgr A* and occupies about only 0.02 pc (Pfahl & Loeb 2004). In this case, their orbital speed around the GC would be $\sim 1000$ km s$^{-1}$ and the birth kick would be unlikely to perturb them significantly from their orbit. This possibility therefore seems unlikely.

The pulsars could originate from a progenitor population within $\sim 1$ pc of the GC. At this distance, the gravitational effect of the black hole is small and the birth process could kick the pulsars out to their current position. A pulsar with a velocity of $100$ km s$^{-1}$ travels 100 pc in 1 Myr, the characteristic age of PSR J1746–2856, and no especially large kick velocity needs to be invoked. In this picture, the pulsars’ proper motion would be in a direction away from the GC.

The pulsars could also originate from a distance of $\sim 200$ pc from the GC inside of which there is known to be significant high mass star formation (Figer et al. 2004). The velocity of the pulsars would be moderate so as to ensure their retention in the GC region. Long term timing of the two pulsars may help to distinguish between these two cases.

6 NO DETECTIONS AT 8.4 GHz

What are the implications of failing to detect any pulsars at 8.4 GHz? Recall that the area surveyed is only a $\pm 15$ pc box around the GC. At the GC distance, the survey region is inside the scattering screen proposed by Lazio & Cordes (1998) and any putative pulsars will therefore suffer from broadening. However, assuming the screen is 130 pc from the GC, the scattering time at 8.4 GHz is then expected to be only $\sim 80$ ms.

Ignoring this effect for now, the sensitivity of the survey is such that a pulsar with a luminosity (at 8.4 GHz) greater than $7$ mJy kpc$^2$ would have been detected. Extrapolating down to 1.4 GHz, assuming a spectral index of $-1.6$, gives a luminosity limit of 125 mJy kpc$^2$. Of the 1160 known pulsars with a flux density measurement at 1.4 GHz, 133 of them (11%) have luminosities in excess of this. Given that there are $\sim 10^3$ active pulsars beaming in our direction and that the known population is not complete even at this luminosity level, at least 0.1% of all pulsars should have luminosities in excess of this value. Therefore, either less than 1000 pulsars beaming in our direction are in the GC region or, as seems more likely, the scattering is severe even at this high frequency. Note that in the inner 1 pc, the luminosity limit is a factor of 3 higher because of the high background temperatures there (see Section 2). This increases the upper limit to the number of detectable pulsars by a similar amount. This is then broadly in line with modelling of the GC population of neutron stars (Cordes & Lazio 1997; Pfahl & Loeb 2004).

7 SUMMARY

We have discovered two pulsars within $0.3^\circ$ of the GC during a targeted survey at 3.1 GHz. Both pulsars have very high DMs and scatter broadening times. It seems unlikely that the survey has penetrated through the scattering screen which surrounds the GC at a distance of $\sim 150$ pc. However, the pulsars are likely to have originated in an excess population associated with the GC and be located within a few hundred pc of it. Both the DM and scattering are then higher than expected in the current electron density models perhaps favouring a filled centre model rather than the existing annular one. No detections were made at 8.4 GHz with the implication being that either there are less than 3000 pulsars beaming in our direction in the inner pc of the Galaxy or that the scattering is more severe than previously thought.

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