MASS ESTIMATORS FOR FLATTENED DISPERSION-SUPPORTED GALAXIES

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ABSTRACT

We investigate the reliability of mass estimators based on the observable velocity dispersion and half-light radius $R_0$ for dispersion-supported galaxies. We show how to extend them to flattened systems and provide simple formulae for the mass within an ellipsoid under the assumption the dark-matter density and the stellar density are stratified on the same self-similar ellipsoids. We demonstrate explicitly that the spherical mass estimators give accurate values for the mass within the half-light ellipsoid, provided $R_0$ is replaced by its “circularized” analog $R_0 \sqrt{1 - \epsilon}$. We provide a mathematical justification for this surprisingly simple and effective workaround. It means, for example, that the mass-to-light ratios are valid not just when the light and dark matter are spherically distributed, but also when they are flattened on ellipsoids of the same constant shape.

**Key words:** dark matter – galaxies: dwarf – galaxies: kinematics and dynamics

1. INTRODUCTION

Accurate estimates of the dark-matter content of dwarf spheroidal galaxies (dSphs) are crucial for furthering our understanding of galaxy formation and structure. Calculating reliable mass estimates has historically been an awkward problem as, with only line of sight (LOS) velocity measurements, the mass profile of a spherical galaxy can only be inferred by making an assumption about the degree of velocity anisotropy, i.e., the ratio of radial to tangential motion.

Through comparisons to solutions of the Jeans equations, it has been shown that the mass contained near the half-light radius of a dispersion-supported galaxy is approximately independent of the velocity anisotropy and the radial profile of the dark and luminous matter and is simply related to the half-light radius $R_0$ and the luminosity-averaged LOS velocity dispersion $\sqrt{\langle \sigma_{\text{los}}^2 \rangle}$. There exist several different forms for these formulae in the literature (Walker et al. 2009; Wolf et al. 2010; Amorisco & Evans 2012; Campbell et al. 2016) that may be summarized as

$$M_{\text{ph}}(<r_s) = \frac{C_s \langle \sigma_{\text{los}}^2 \rangle R_0}{G}$$  \hspace{1cm} (1)

where $M_{\text{ph}}(<r_s)$ is the mass contained within a sphere of radius $r_s$ and $G$ the familiar gravitational constant. $C_s$ is a constant that depends on the choice of radius $r_s$. Walker et al. (2009) proposed that if $r_s = R_0$, then $C_s = 2.5$ based on a simple example of the stellar distribution following a Plummer profile and the dark matter following a cored isothermal profile, although this was validated through fuller testing. Wolf et al. (2010) demonstrated that for $r_s \approx \frac{2}{3} R_0$ (approximately the 3D spherical half-light radius for a range of observationally motivated profiles) that $C_s = 4$ reproduced the results from full Jeans analyses and was also shown to be mathematically true under the assumption of a near-flat velocity dispersion profile.

Although spherical mass estimators have proved useful for understanding dSphs, they cannot give the full picture as they do not consider the fundamentally aspherical shape of these galaxies. Our aim in this Letter is to find mass estimators equivalent to Equation (1) applicable to flattened systems. We begin by inspecting the validity of the spherical mass estimators and go on to investigate the applicability of the estimator when considering flattened systems in which the dark and light matter are stratified on the same self-similar ellipsoids. We give formulae similar to Equation (1) that may be used when the 3D shape of the system is known. By marginalizing over prior assumptions on the intrinsic shape and alignment, we show how the mass can be estimated when the intrinsic shape and alignment are not known.

2. SPHERICAL MASS ESTIMATORS

For a spherical stellar luminosity density $j_\ast(r)$ with a constant mass-to-light ratio in a spherical mass density $\rho_{\text{DM}}(r)$ with mass profile $M(r)$ sourcing potential $\Phi(r)$, the potential energy can be written in terms of the surface brightness $S(R)$ as

$$W = \frac{1}{2} \int dV \, j_\ast(r) \Phi(r) = 4\pi G \int_0^\infty dr \, I(r) M'(r),$$

where

$$I(r) = \int_r^\infty dr \, j_\ast(r) = -\frac{1}{\pi} \int_r^\infty dR \, (R^2 - r^2)^{1/2} \frac{dS}{dR}.$$  \hspace{1cm} (3)

From the virial theorem, we know that the LOS velocity dispersion is related to the total luminosity $L$ by

$$\langle \sigma_{\text{los}}^2 \rangle = -W/3L,$$  \hspace{1cm} (4)

which gives the constant $C_s$ as

$$C_s = \frac{1}{R_0} \left[ \int_0^{\infty} dr \, r^2 \rho_{\text{DM}}(r) \right]^{1/2} \int_0^{\infty} dr \, r^2 J(r) \rho_{\text{DM}}(r)^{-1}.$$  \hspace{1cm} (4)

where $J(r) = (4\pi/3) L I(r)$. The constant $C_s$ depends only on the profile of the halo model $\rho_{\text{DM}}$ and the surface brightness profile $I(r)$.

We use this to test the validity of the spherical mass estimator. In Figure 1, we show the result of Equation (4) computed numerically for two models with differing ratios of dark to stellar scale-lengths ($\eta_{\text{DM}}/R_0$). They are an NFW dark-matter profile $\rho_{\text{DM}}(r) \propto r^{-1}(1 + r/\eta_{\text{DM}})^{-2}$ and a cored isothermal profile of the form

$$\rho_{\text{DM}}(r) = \frac{\eta_{\text{DM}}^2}{4\pi} \frac{3\eta_{\text{DM}}^2 + r^2}{(\eta_{\text{DM}}^2 + r^2)^2}.$$  \hspace{1cm} (5)
The stellar tracer profile follows a Plummer law \( \rho_i(r) \propto (1 + (r/r_5)^2)^{-5/2} \) for which \( R_5 = r_5 \). The constant \( C_i \) is computed at the two radii recommended by Walker et al. (2009) and Wolf et al. (2010). The constants given by these two authors are shown with horizontal lines along with the uncertainty found by Campbell et al. (2016) from inspecting cosmological hydrodynamical simulations. The variation of \( C_i \) with respect to \( \eta_{DM}/R_h \) is smallest for the NFW profile and is consistent with the bracket found by Campbell et al. (2016). In the cored isothermal profile with \( \eta_{DM}/R_h \approx 1 \), both estimators perform well. However, again, as \( \eta_{DM}/R_h \) is increased, \( C_i \) deviates significantly and so the estimators perform poorly for \( \eta_{DM}/R_h > 2 \).

We now explore how the mass estimators perform as the parameters of a double power-law dark-matter density profile are altered. We use a fixed Plummer profile for the stars with a sech truncation at 10\( R_h \). In Figure 2, we show the mass profiles of different dark-matter profiles that all produce the same luminosity-averaged LOS velocity dispersion. The default parameters are those of an NFW profile with \( R_{DM}/R_h = 1 \) and a sech truncation at 10\( R_{DM} \). We alter the outer slope \( \beta \), inner slope \( \gamma \), and the ratio \( R_{DM}/R_h \). We find that when varying the inner and outer slopes the pinch point where the mass is the same for all profiles is around \( 4/3 \) the radius recommended by Wolf et al. (2010). Varying \( R_{DM}/R_h \) produces a pinch point further out. This helps explain why mass estimators derived for use on realistic halos with \( R_{DM} > R_h \) can constrain the mass at larger radii (e.g., Amorisco & Evans 2012; Campbell et al. 2016).

3. FLATTENED MASS ESTIMATORS

We now turn to adapting the spherical mass estimators for application to flattened systems. We work with models with both the dark and stellar density stratified on the same concentric self-similar ellipsoids labeled with the coordinate \( m = x^2/a^2 + y^2/b^2 + z^2/c^2 \) with \( a > b > c \). The axis ratios of the ellipsoids are \( p = b/a \) and \( q = c/a \). We view the model along the spherical polar unit vector defined by the angles \( (\theta, \varphi) \), where \( \theta \) is the co-latitudinal angle and \( \varphi \) the azimuthal angle defined with respect to a Cartesian coordinate system aligned with the principal axes (see Figure 3). When oblate and prolate spheroids are viewed “face-on,” they appear round. The spherical mass estimator underestimates (over-estimates) the mass within a sphere for the oblate (prolate) case, as mass is added to (removed from) the sphere. Similarly, the formulae give (smaller) under- and overestimates for the mass within the corresponding ellipsoid. We seek an appropriate modification to Equation (1) that is applicable to flattened systems, namely,

\[
M_{ell}(<m_c) = C_f f_r \langle \sigma^2_{los} \rangle f_r R_h / G,
\]

where \( M_{ell}(<m_c) \) is the mass within an ellipsoid (that is the same shape as the equipotential contours) with major axis length \( r_s \). We imagine creating an elliptoidal model by deforming a spherical model that obeys the spherical mass estimator formulae outlined in the previous section. The total mass is conserved if \( abc = 1 \) and the mass within an ellipsoid of major-axis length \( r_s \) is identical to the mass within a sphere of radius \( m_c = r_s/a = r_s/(pq)^{1/3} \). However, to estimate this parent spherical mass model, we must relate the observed LOS velocity dispersion to the spherical velocity dispersion and the observed half-light major-axis length to the intrinsic major-axis length of the considered ellipsoid. Assuming the total velocity dispersion (the average of the dispersions along the principal axes) is conserved as we deform the model\(^1\), the factor \( f_\sigma \) accounts for the relationship between the LOS velocity dispersion and the total dispersion of the elliptoidal model. The factor \( f_r \) accounts for the relationship between the observed major-axis length and the intrinsic major-axis length of the equivalent ellipsoid (and that of the parent spherical model).

3.1. Velocity Scaling

For triaxial systems, the velocity scaling \( f_\sigma = \langle \sigma^2_{los} \rangle / \langle \sigma^2_{los} \rangle \) is given by

\[
f_\sigma = \frac{1}{3} \frac{1 + r_{cx} + r_{cz}}{r_{cx} \sin^2 \theta \cos^2 \varphi + r_{cz} \sin^2 \varphi \sin^2 \varphi}
\]

where

\[
r_{ij} = \langle \sigma^2_{ij} \rangle / \langle \sigma^2_{ij} \rangle = W_i/W_j \quad (\text{no sum}).
\)

For dSphs in which the stellar and dark-matter density profiles are stratified on the same self-similar ellipsoids, \( r_{ij} \) depends only on the shape of the ellipsoids (Roberts 1962; Binney & Tremaine 2008). That is to say, it is independent of the “radial” density profile of the light and dark matter. Therefore, \( f_\sigma \) is a

\(^1\) To leading order in the flattening, the ratio of the total dispersion of the flattened model to the spherical model with the same mass is \( \langle \sigma^2_{los} \rangle / \langle \sigma^2_{los} \rangle \approx 1 - \frac{1}{(1 - p)^2(1 - q) + (1 - q)^2} \).
related to the intrinsic coordinate system by

\begin{align*}
W_{ij} &= f_{ij}(\theta, \varphi, p, q).
\end{align*}

Expressions for \( W_{ij} \) are given in Table 2.2 of Binney & Tremaine (2008).

3.2. Radial Scaling

We decompose the radial scaling into two components, \( f_r = f_1 f_2 \). \( f_2 \) describes the relationship between the ellipsoidal major-axis length and the parent spherical radius \( r_h \) (as described above) \( f_2 = (pq)^{1/3} \). The other factor \( f_1 \) gives the relationship between the observed major-axis length of the half-light ellipse \( R_h \) and the intrinsic major-axis length of the corresponding ellipsoid \( r_{maj} \). In the spherical case, these quantities are equal. In the ellipsoidal case, the relationship between these quantities depends on the viewing angles and the intrinsic shape \( f_1 = f_1(\theta, \varphi, p, q) \). We approximate \( f_1 \) by the relationship between the major-axis length of an ellipsoid and the major-axis length of its projected ellipse. This neglects any subtleties related to the extended nature of the true density distribution. However, if the 3D stellar light profile falls off sufficiently rapidly, then our relationship is a good approximation.

To derive our approximation for \( f_1 \), we use a coordinate system \((x', y', z')\) related to the intrinsic coordinate system by (see Figure 3)

\begin{align*}
&x = -x' \sin \varphi - y' \cos \theta \cos \varphi + z' \sin \theta \cos \varphi \\
y = x' \cos \varphi - y' \cos \theta \sin \varphi + z' \sin \theta \sin \varphi \\
z = y' \sin \theta + z' \cos \theta.
\end{align*}

We consider the set of points where the ellipsoidal surface is tangential to \( \xi' \) which results in a rotated ellipse in the \((x', y')\) plane. We diagonalize the resultant quadratic surface to find the major axis length \( R_h \) as

\begin{equation}
\begin{aligned}
f_1^{-2} &= \left( R_h / r_{maj} \right)^2 = 2C / (A - \sqrt{B}),
\end{aligned}
\end{equation}

where

\begin{equation}
\begin{aligned}
A &= (1 - q^2) \cos^2 \vartheta + (1 - p^2) \sin^2 \vartheta \sin^2 \varphi + p^2 + q^2, \\
B &= [(1 - q^2) \cos^2 \vartheta - (1 - p^2) \sin^2 \vartheta \sin^2 \varphi - p^2 + q^2]^2 \\
&+ 4(1 - p^2)(1 - q^2) \sin^2 \vartheta \cos^2 \vartheta \sin^2 \varphi, \\
C &= p^2 \cos^2 \vartheta + q^2 \sin^2 \vartheta (p^2 \cos^2 \varphi + \sin^2 \varphi).
\end{aligned}
\end{equation}

As given in Weijmans et al. (2014), the observed ellipticity \( \epsilon \) satisfies 

\begin{equation}
(1 - \epsilon)^2 = (A - \sqrt{B}) / (A + \sqrt{B}),
\end{equation}

function of \( p, q \), and the viewing angles: \( f_{ij} = f_{ij}(\theta, \varphi, p, q) \).

\begin{figure}[h]
\centering
\includegraphics[width=\textwidth]{sph_mass_est.png}
\caption{Spherical mass profile for a stellar Plummer profile embedded in a double power-law dark-matter halo with varying outer slope \( \beta \), inner slope \( \gamma \), and scale radius \( r_{DM} \). All models have the same luminosity-averaged LOS velocity dispersion and half-light radius \( R_h \). The masses are normalized with respect to the Wolf mass estimate. The default parameters are \( \gamma = 1 \), \( \beta = 3 \) (NFW), and \( r_h / R_h = 1 \). The black points show the results of two mass estimators and the vertical dashed line shows the point of minimum variance in the logarithm of the mass for each set of curves. The spheroidal mass estimates using the mass estimator proposed in this Letter are given for an edge-on oblate \( (q = 0.6, \text{ orange triangles}) \) and edge-on prolate model \( (p = q = 0.6, \text{ pink triangles}) \).}
\end{figure}

\begin{figure}[h]
\centering
\includegraphics[width=\textwidth]{radial_scaling.png}
\caption{Relationship between the observed half-light radius and the radius used in our mass estimator formula. An ellipsoid is observed at spherical polar angles \((\theta, \varphi)\) with respect to its intrinsic Cartesian coordinates \((x, y, z)\) aligned with the principal axes. The resulting projection is an ellipse (shown below) with major-axis length \( R_h \) that lies in the \((x', y')\) plane of the observed Cartesian coordinate system \((x', y', z')\). Above the ellipsoid, we show the sphere with the equivalent volume as the ellipsoid. The major axis of the ellipse is related to the major axis of the ellipsoid by the factor \( f_{1} \), which is shown in the lower two panels for a prolate spheroid with \( p = q = 0.5 \) (left) and an ellipsoid with axis ratios \( p = 0.85 \) and \( q = 0.5 \) (right).}
\end{figure}
For an oblate spheroid ($p=1$), Equation (10) simplifies to $R_h = r_{maj}$. For a prolate spheroid $p = q$, so we find

$$f_1^{-2} = \cos^2 \vartheta + \sin^2 \vartheta (q^2 \cos^2 \varphi + \sin^2 \varphi). \quad (12)$$

In Figure 3, we show the major axis length for a prolate figure and a triaxial figure as a function of the viewing angle.

The ellipsoidal half-light radius $m_h$ is well approximated by $\frac{4}{3} f_r R_h$, which should be compared to the radius of $\frac{4}{3} R_h \sqrt{1 - \epsilon}$ that is empirically used (e.g., Koposov et al. 2015; Sanders et al. 2016) as $R_h \sqrt{1 - \epsilon}$ approximately reproduces the circularly averaged half-light radius of the dSph.$^2$

### 3.3. Near-spherical Limits

Using Equation (10), we can find the modification factor $f_{sr} f_r$ for the simple cases of viewing down the principal axes of a near-spherical triaxial ellipsoid and compare to the alternative factor $\sqrt{1 - \epsilon}$. When viewing down the major axis ($\vartheta = \pi/2$, $\varphi = 0$), we find

$$f_{sr} f_r \approx 1 + \frac{2}{3}(1 - p) - \frac{1}{3}(1 - q). \quad (13)$$

The observed ellipticity $\epsilon = 1 - p/q$ so the circularized factor $\sqrt{1 - \epsilon} \approx 1 + \frac{1}{2}(1 - p) - \frac{1}{4}(1 - q)$, which is a close approximation to our factor $f_{sr} f_r$. Similarly, for viewing down the intermediate axis, we find

$$f_{sr} f_r \approx 1 + \frac{1}{2}(1 - p) - \frac{1}{3}(1 - q), \quad (14)$$

while $\sqrt{1 - \epsilon} \approx \sqrt{q} \approx 1 - \frac{1}{2}(1 - q)$. Finally, viewing down the minor axis, we find

$$f_{sr} f_r \approx 1 - \frac{1}{2}(1 - p) + \frac{1}{2}(1 - q), \quad (15)$$

while $\sqrt{1 - \epsilon} \approx \sqrt{p} \approx 1 - \frac{1}{4}(1 - p)$. We note that the flattening in the LOS direction (e.g., $p$ in the intermediate axis case) has a smaller contribution to the factor $f_{sr} f_r$. This demonstrates that the simple factor $\sqrt{1 - \epsilon}$ goes a long way to account for the velocity and radial scalings we propose.

### 3.4. Results

In Figure 2, we show the mass estimates using our formulae for an oblate and prolate model viewed edge-on. The models have the same ellipsoidal mass profile as the spherical model shown. The factors we have introduced correctly deproject the observed quantities producing an unbiased mass estimate. Figure 4 shows the constant in the half-light ellipsoid mass estimator (Equation (6)) for three models of flattened Plummer profiles embedded in equivalently flattened NFW halos with $m_{200}/m_\odot = 5$. We show an oblate, prolate, and triaxial $p = \frac{1}{2}(1 + q)$ model. Simply using the spherical mass estimator with $R_h$ underestimates/overestimates the ellipsoidal mass for the oblate/prolate case viewed face-on (down the minor/major axis). Similarly, the edge-on case (major for oblate, minor for prolate) produces overestimates of the mass for both oblate and prolate models. For the triaxial model the spherical mass estimator produces an overestimate when viewing down the major axis and (for this particular case) is largely unbiased when viewing down the minor axis. The results using the correction factors $f_r$ and $f_p$ are unbiased estimates of the mass within the ellipsoid $m = \frac{4}{3} m_\odot$ and using the spherical mass estimator with the “circularized” radius $R_h \sqrt{1 - \epsilon}$ produces very similar results to the corrected version. This echoes a result in Sanders et al. (2016), who demonstrated that the correction to the $D$-factor (important for interpreting dark-matter decay signals) is almost independent of the flattening for edge-on systems. The near-spherical expansions of Section 3.3 are also shown, which replicate the trends over the full $q$ range.

Our proposed modifications correctly reproduce the mass within ellipsoids. However, this relies on knowing the intrinsic shape and alignment of the dSph. Such information is not accessible, but we can put priors on possible models that reproduce the observables. We choose to put priors on the triaxiality $T = (1 - p^2)/(1 - q^2)$, flattening $q$, and the viewing angles ($\vartheta$, $\varphi$). We consider three priors:

1. Flat prior—$T \sim U(0, 1)$, $q \sim U(0.05, 1)$, $\cos \vartheta \sim U(0, 1)$, $\varphi \sim U(0, \pi/2)$.
2. Major-axis prior—$T \sim U(0, 1)$, $q \sim U(0.05, 1)$, $\vartheta \sim N(\pi/2, 0.1$ rad), $\varphi \sim U(0, 0.1$ rad).
distribution of $f_{\epsilon}/\sqrt{1-\epsilon}$ for each prior assumption (i.e., the ratio of the size of the ellipsoid to the $R_h\sqrt{1-\epsilon}$ approximation). For the uniform prior, this ratio is unity (within $\sim10\%–20\%$ so the “size” of the dSphs are well approximated by $R_h\sqrt{1-\epsilon}$). For the other two priors, the ratio increases with ellipticity as the intrinsic ellipsoids are on average more elongated along the LOS so larger than $R_h\sqrt{1-\epsilon}$.

We have demonstrated that the mass within the half-light ellipsoid can be accurately estimated using the spherical mass estimator formulae. Although we do not know the shape or orientation of this half-light ellipsoid, we can say with confidence the mass within it. Therefore, we can accurately estimate the mass-to-light ratio using the mass within the half-light ellipsoid and half the total luminosity $L$. We conclude that using the spherical mass estimators (Walker et al. 2009; Wolf et al. 2010) with the “circularized” half-light radius produces accurate estimates of the mass-to-light ratio of dSphs, irrespective of flattening, provided the light and dark matter are stratified on the same self-similar concentric ellipsoids.

4. CONCLUSIONS

This Letter has answered the question: how should the mass of a flattened, dispersion-supported galaxy like a dwarf spheroidal be estimated? If the galaxy were spherical, then the answer is well established. Accurate mass estimators depending on the observable half-light radius and the velocity dispersion of the stars have been devised by a number of investigators (Walker et al. 2009; Wolf et al. 2010; Amorisco & Evans 2012; Campbell et al. 2016).

We have shown how to modify the spherical mass estimators so that they work for flattened systems in which the light and dark matter are stratified on the same self-similar ellipsoids. This represents a limiting case as simulations indicate the dark-matter distribution is in fact rounder than the light (Abadi et al. 2010; Zemp et al. 2012) due to baryonic feedback effects, particularly for the more massive dSphs. The modifications require knowledge of the intrinsic shape and alignment of the triaxial figure and reproduce the mass within ellipsoids by deprojecting the half-light radius and LOS velocity dispersion. The resulting mass estimates are independent of details of the radial profile and are as accurate as the corresponding spherical formulae.

This would be of little use if we require knowledge of intrinsic properties. However, we have also shown that, when averaging over triaxial configurations that are consistent with the observed ellipticity $\epsilon$, major-axis half-light length $R_h$, and LOS velocity dispersion, the mass within the half-light ellipsoid is well approximated by the spherical mass estimate using the “circularized” half-light radius of $R_h\sqrt{1-\epsilon}$. The scatter in the estimate increases with ellipticity but is only $10\%–20\%$ for $\epsilon \sim 0.6$. In turn, this observation implies that mass-to-light ratios using spherical estimators, together with a luminosity of $L_{1/2} = L/2$, are accurate and insensitive to the flattening of the dSph. This therefore provides a surprisingly simple, flexible, and effective way to account for the effects of flattening.

REFERENCES

Abadi, M. G., Navarro, J. F., Fardal, M., Babul, A., & Steinmetz, M. 2010, MNRAS, 407, 435
Amorisco, N. C., & Evans, N. W. 2012, MNRAS, 419, 184

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3 We note that a similar observation was made by Laporte et al. (2013), who found that the variations of $R_h$ and $(\sigma^2_{\text{rms}})$ with triaxiality compensated each other to give an unbiased mass estimate.
Barber, C., Starkenburg, E., Navarro, J. F., & McConnachie, A. W. 2015, *MNRAS*, 447, 1112
Binney, J., & Tremaine, S. 2008, Galactic Dynamics (2nd ed.; Princeton, NJ: Princeton Univ. Press)
Campbell, D. J. R., Frenk, C. S., Jenkins, A., et al. 2016, arXiv:1603.04443
Foreman-Mackey, D., Conley, A., Meierfjngen Farr, W., et al. 2013, emcee: The MCMC Hammer, Astrophysics Source Code Library, ascl:1303.002
Koposov, S. E., Belokurov, V., Torrealba, G., & Evans, N. W. 2015, *ApJ*, 805, 130
Laporte, C. F. P., Walker, M. G., & Peñarrubia, J. 2013, *MNRAS*, 433, L54
Roberts, P. H. 1962, *ApJ*, 136, 1108
Sánchez-Janssen, R., Ferrarese, L., MacArthur, L. A., et al. 2016, *ApJ*, 820, 69
Sanders, J. L., Evans, N. W., Geringer-Sameth, A., & Dehnen, W. 2016, *PhRvD*, 94, 063521
Walker, M. G., Mateo, M., Olszewski, E. W., et al. 2009, *ApJ*, 704, 1274
Weijmans, A.-M., de Zeeuw, P. T., Emsellem, E., et al. 2014, *MNRAS*, 444, 3340
Wolf, J., Martinez, G. D., Bullock, J. S., et al. 2010, *MNRAS*, 406, 1220
Zemp, M., Gnedin, O. Y., Gnedin, N. Y., & Kravtsov, A. V. 2012, *ApJ*, 748, 54