Optical/Ultraviolet Continuum Emission Theory in Radio Quiet Quasars and Active Galactic Nuclei

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Abstract. Accretion disk models still do not provide a satisfactory explanation of the optical/ultraviolet continuum observed in Seyferts and quasars. Substantial theoretical progress has been made in understanding one aspect of the problem: the dearth of spectral features at the Lyman limit. Promising solutions have also been proposed to explain the surprising observations of large polarization in the Lyman continuum observed in some sources. I review the recent progress in this field, and try to point out future research directions which would be fruitful in trying to obtain a complete, self-consistent model of the continuum emitting regions.

INTRODUCTION

Thermal emission from accretion disks has long been the standard paradigm for the optical/ultraviolet (OUV) continuum in quasars and active galactic nuclei (AGN). However, as noted in the talk by Koratkar in this conference [1], the simplest, bare, geometrically thin, optically thick accretion disk models face a number of severe problems when confronted with observations. These include the observed absence of absorption or emission edges at the Lyman limit in most sources [2,3], the continuum polarization [4,5], the continuum spectral energy distribution [6], phased optical/ultraviolet variability [7–9], and microlensing constraints in the Einstein Cross [10,11]. In addition, there are issues around connecting the ultraviolet continuum to photoionization models of the emission line regions [12], and how the OUV continuum is related to the soft and hard X-rays observed from these sources. As Koratkar emphasizes in her talk, a theorist’s job is to design a predictive model which explains all the observed phenomenology simultaneously, not just in a piece-meal fashion. Such a model does not yet exist, although progress is definitely being made. In this paper I will review some of the progress, concentrating in particular on the Lyman edge problem because it sheds light on many of the other issues.
THE LYMAN EDGE PROBLEM

The Lyman edge problem arose from early calculations of AGN accretion disk spectra by Kolykhalov & Sunyaev (1984) [13]. These authors adopted the reasonable approach of taking existing stellar atmosphere libraries at that time [14–16], assigning a spectrum to each radius of an alpha disk [17] by using an atmosphere model which had the same effective temperature and surface gravity, and then integrating over the disk to produce the resulting spectrum. All the disk models they explored had very large absorption edges at the Lyman limit, and they noted that these disagreed with observations of bright quasars at the time. This discrepancy motivated much of the later theoretical and observational work on the Lyman edge region of AGN spectra, but it is important to note the limitations of these early calculations. The stellar atmosphere libraries which were used completely neglected non-LTE effects. These can be very important in the Lyman edge region because the large hydrogen photoionization opacity generally means the Lyman continuum originates high in the atmosphere where densities are low and radiative transition rates can dominate collisional transition rates. Even more important, the atmospheres which exist in those libraries generally have high densities compared to what is possible in high luminosity accretion disks. This limited the range of parameter space which Kolykhalov & Sunyaev could explore to low alpha parameter ($\alpha < 10^{-2}$), high photospheric density models. Because the neutral fraction of hydrogen becomes larger at high densities, this resulted in very large Lyman edge absorption opacity, thereby producing very large absorption edges in these disk models.

As discussed in more detail by Koratkar [1], spectral features near the Lyman limit (whether in emission or absorption) are very rare in quasars, except in cases where there is clearly evidence for intervening absorbers [2,3]. There are cases of weak features which are consistent with being intrinsic to the quasar, but only in $\sim$ 10% of quasars studied [3]. Intervening absorption by the intergalactic medium is less of a problem for low redshift Seyferts, but one must then contend with the absorption due to our own Galaxy. Hopkins Ultraviolet Telescope data generally show no features at the intrinsic Lyman limit of low redshift type 1 Seyferts [18], including NGC 5548. This is particularly disturbing as NGC 5548 has a broad iron Kα line, which is strong evidence for a relativistic accretion disk [19,20]. There is some absorption at slightly shorter wavelengths, but this could conceivably be due to high $n$ Lyman series lines from the Galaxy. Zheng et al. (1995) [21] find that an intrinsic Lyman edge provides a good fit to the data for the type 1 Seyfert Mrk 335. More data and careful modeling is clearly needed.

The Hubble Space Telescope (HST) composite quasar spectrum of Zheng et al. (1997) [22] shows at most only a very shallow absorption feature at the Lyman limit, but it does show a spectral break near there. The spectrum shortward of the Lyman limit can be fit with a power law, and extension of that power law lines up approximately with the ROSAT composite spectrum of Laor et al. (1997) [23]. This spectral break near the Lyman limit actually causes problems in another
area of AGN phenomenology, that of the emission line regions. If the line emitting gas sees the same photoionizing continuum as we observe in the composite, then there are too few photons to produce the observed He II emission lines [12]. It is important to remember that Zheng et al. had to correct for the Lyman α forest and Lyman continuum absorption by intervening material. In addition, they did not correct for internal extinction, which could be very important in the ultraviolet. (This may nevertheless be okay given the lack of any obvious 2200Å feature.) The composite spectrum may indicate that many quasars do have absorption at the Lyman limit combined with e.g. Comptonization, as the authors suggest. It would be nice to see this spectral energy distribution in an individual source, however.

There have been essentially three theoretical approaches to solving the Lyman edge problem within the standard optically thick, geometrically thin accretion disk paradigm: more sophisticated stellar atmosphere modeling, Comptonization by a hot corona, and relativistic smearing.

**Stellar Atmosphere Effects**

Low densities in the disk photosphere will reduce the ratio of Lyman continuum absorption opacity to scattering opacity, and thereby decrease the flux difference at the absorption edge, or even drive it into emission [24]. Ab initio radiative transfer calculations confirm this behavior [25–28], showing that as the maximum effective temperature and/or the ratio of luminosity to Eddington luminosity increase, Lyman absorption edges become weaker, then disappear, become moderate emission edges, and finally become weak emission edges as the hydrogen becomes fully ionized and recombination rates are low. Non-LTE effects, where the ionization states of hydrogen are determined at least partially by the radiation field rather than just by collisions, produce similar results [29,26,30,27]. (They can also greatly enhance the He II Lyman continuum [27], which may help photoionization models of the emission line regions [12].) To completely get rid of the edge in a given atmosphere using these effects would require fine-tuning, but it is clear that the overall edge produced from an integration of diverse atmospheres at different disk radii, some with absorption edges and some with emission edges, might generically produce a small net effect.

Most of the radiative transfer calculations in AGN accretion disk theory have been fairly crude or incomplete, and it is only recently that modern, state of the art, stellar atmosphere codes have been applied to the problem. A crucial effect that has generally been neglected in the past because of its complexity is metal line blanketing. Hubeny & Hubeny have shown at this conference [31] that metal lines can produce much stronger spectral features in an individual atmosphere in the Lyman edge region than the Lyman edge itself. In fact, judging from their plots, an observer would be hard-pressed to see any Lyman edge feature, even if they could observe the atmosphere directly without the effects of Comptonization and relativistic smearing discussed below.
All these stellar atmosphere models neglect the effects of external illumination of the atmosphere. AGN and quasars emit X-rays, and in Seyferts, X-rays carry a substantial fraction of the bolometric luminosity. Indeed, in the most extreme form of the currently popular two phase disk/corona accretion model [32], all the accretion power is initially converted into heat in the corona above the disk. Disk heating would therefore take place externally, rather than internally. Real sources probably have a combination of internal and external heating, such as in the patchy corona model of Haardt, Maraschi, & Ghisellini (1994) [33]. Sincell & Krolik (1997a) [34] have done the best job so far of calculating the resulting spectrum from an X-ray illuminated disk, starting with the extreme assumption of total external heating. The lack of internal, viscous heating results in a much lower disk scale height and therefore a much larger photospheric density. As a result, they found very substantial Lyman edges in emission, in disagreement with observation. This calculation would be worth doing even better, however. The X-rays were incorporated just through a depth-dependent heating function, and were therefore not included as part and parcel of the radiation field in the radiative transfer and atmosphere structure calculation. In addition, the calculation was done in LTE, and therefore X-ray and UV photoionization effects on the ground state population of hydrogen were neglected. Even so, the trend of large edges produced with high photosphere densities may be robust, and may spell trouble for the simplest two-phase models.

It is perhaps useful to introduce a reality check here by comparing how well stellar atmosphere models do in reproducing the Lyman edges in real stars, which we know are optically thick! The B2 II star $\epsilon$ CMa exhibits a substantial Lyman absorption edge, but much less than predicted by both LTE and non-LTE models with line blanketing [35]. In fact, the Lyman continuum in this star exceeds that predicted by the stellar atmosphere models by a factor $\sim 30!$ It is interesting that one proposed solution of this problem involves external illumination of the star's atmosphere by X-rays [36]. Stellar atmosphere models also fail to reproduce the spectrum of the B1 II-III star $\beta$ CMa in all wavelength regions [37].

**Comptonization**

If X-rays are produced by Comptonization in a hot corona above the disk, then the optical/ultraviolet spectrum will itself be modified in the process. Comptonization can substantially reduce a Lyman absorption edge by scattering low energy photons above the edge, thereby filling it up [38–40]. Downscattering by a hot corona can also reduce the magnitude of an emission edge discontinuity, although not as dramatically for the same coronal optical depth and temperature because of the tendency for upscattering in a hot corona [40]. In either case, a spectral break in the continuum slope is generically produced for modest optical depths, and this has been used to fit the HST composite spectrum [22].

Antonucci (1992) [4] has criticized this solution to the Lyman edge problem on the grounds that the modest scattering depths required would impart substantial
polarization everywhere in the optical/ultraviolet, which is not observed. In fact Compton scattering generally produces less polarization in a featureless continuum source than Thomson scattering, but nevertheless substantial polarization does result if an edge is successfully smeared out [40]. It is likely that a Comptonizing corona is powered through magnetic fields, and if these are anywhere near equipartition strength with the radiation field, then Faraday rotation can greatly reduce the polarization (see below).

Relativistic Doppler Shifts and Gravitational Redshifts

Many authors have invoked the highly relativistic orbital velocities and varying gravitational redshift with distance from the hole to smear out the Lyman edge (e.g. [29,41,42,39,25,26,43]). Not surprisingly, Kerr holes are more effective at smearing edges than Schwarzschild holes because of the deeper potential wells. In addition, edge-on disks are more effective than face-on disks. However, it should be born in mind that type I objects are probably viewed more face-on than edge-on, at least according to the unified model [44] and the inferences from the iron Kα lines [20]. It is interesting to note that the paper by Kolykhalov & Sunyaev (1984) [13], which first articulated the Lyman edge problem as discussed above, did in fact include relativistic smearing by a Schwarzschild hole. This failed to remove the huge Lyman discontinuities which were produced by their dense atmosphere models.

Agol, Hubeny, & Blaes have presented results at this conference [45] of line-blanketed stellar atmosphere models (cf. [31]) folded through the relativistic transfer function of Agol (1997) [46]. The individual atmosphere models are not yet fully self-consistent, and more parameter space needs to be explored, but the results so far indicate very weak Lyman edges.

Polarization

A long-standing criticism of geometrically thin, optically thick accretion disk models is that the emerging radiation should be polarized. Chandrasekhar (1960) [47] showed that an optically thick, pure electron scattering atmosphere produces a polarization of up to 11.7% parallel to the atmosphere plane. Real accretion disks, however, also have significant absorption opacity [48] and may not be exactly planar [49]. In addition, if photospheric magnetic fields are anywhere near equipartition strength, then the optical radiation field could be completely depolarized by Faraday rotation [50–56]. The observed optical polarization may be due to scattering by material further out from the central engine (e.g. [57]).

The Lyman limit region has again turned out to be extremely interesting with regard to polarization. As reviewed in more detail by Koratkar in this conference [1], some quasars show steep rises in polarization shortward of the Lyman limit [58,59], in one case to \( \sim 20\% \)! The samples are not statistically homogeneous, but there appears to be an association between the presence of a partial Lyman
absorption edge and significant polarization shortward of the edge, although this is
not always the case [60].

Substantial anisotropy in the radiation field (limb darkening) at short wave-
lengths can be produced in a cool disk atmosphere because of the steep increase in
the source function as the temperature rises with depth. This effect can produce a
sharp rise in polarization shortward of the Lyman limit which qualitatively agrees
with the observations [61], although quantitative agreement with the data has not
yet been obtained. Relativistic smearing of the Lyman edge also blueshifts and
reduces the rise in polarization [46,62]. It is possible to smear out the edge in total
flux and maintain a steep polarization rise if the optically thick portion of the disk
does not extend all the way down to the innermost stable circular orbit, e.g. if there
is an advection dominated accretion flow in the inner region. Comptonization also
tends to smear out a steeply rising polarization emerging from an underlying disk
photosphere, although again, there are regions of parameter space where the edge
is substantially reduced in total flux while maintaining a steep polarization rise
[40].

Shields, Wobus, & Husfeld (1997) [62] have recently come up with an interesting
model which, while \textit{ad hoc}, produces excellent quantitative agreement with the ob-
servations. They assume that each annulus of the disk locally produces a spectrum
with a large Lyman edge in absorption. In addition, they assume that the local
polarization longward of the edge is near zero, while shortward of the edge it is
\textit{arbitrarily} large. After folding through the relativistic transfer function, this model
can fit the data with few free parameters. Shields et al. suggest that this toy model
might be produced by an optically thin, ionized region of the disk which emits Ly-
man continuum photons which are then polarized by scattering. This ionized gas
might be produced by photoionization by X-rays from a corona, or perhaps be an
optically thin, inner region of an otherwise optically thick disk. We have also found
recently that Comptonization of an absorption edge reproduces the features of their
toy model quite naturally [40]. The polarization shortward of the edge is negative,
in the sense that the plane of polarization is parallel to the disk axis, in contrast to
the polarization produced by limb darkening invoked by [61]. However, it remains
to be seen whether a spectrum produced by a self-consistent atmosphere/corona
folded through the relativistic transfer function can truly reproduce the data.

\textbf{ALTERNATIVES TO THE OPTICALLY THICK
ACCRETION DISK PARADIGM}

The problems with the standard accretion disk model have led some authors to
explore alternatives based on optically thin thermal emission (e.g. [63–68]). The
fact that single temperature thermal bremsstrahlung radiation provides a simple
quantitative fit to the observed continuum spectral energy distributions should
perhaps not be dismissed lightly. However, optically thin models also generally
produce very significant bound-free and bound-bound emission, and therefore suffer
similar problems. Again, Comptonization and relativistic smearing can be invoked to get rid of these spectral features. The broad iron Kα lines [20] and X-ray reflection humps [69] observed in Seyferts are a strong argument for the presence of slabs of optically thick material, such as would be present in the standard accretion disk paradigm. Hybrid models of magnetically confined clouds above an accretion disk have also been considered (e.g. [68]).

It may very well be that a multiphase medium will be necessary to explain the entire spectral energy distribution from optical to hard X-rays of Seyferts and quasars. Preliminary work by Magdziarz et al. (1997) [70] invokes such a mixture of a cold disk, hot X-ray emitting plasma, and an intermediate temperature, optically thick, Comptonizing phase to explain the OUV, soft X-ray excess, and hard X-rays in NGC 5548. More simultaneous, multiwavelength variability campaigns are necessary to provide us with the data required to resolve the geometry of the continuum emission regions and to sort out the energetics of the reprocessing.

**CONCLUSIONS**

In the particular problem of the absence of Lyman edge features in quasars, theory and observation appear to be converging, and I believe that realistic accretion disk models will have little trouble in explaining the absence of such edges in total flux in the majority of objects. Whether this will extend to Seyfert galaxies remains to be seen, and requires additional spectroscopy data in the Lyman limit region of sources at moderate redshifts (hopefully which also exhibit broad iron Kα lines) as well as more theoretical calculations of the effects of external illumination by X-rays. There are promising solutions proposed for the polarization observations around the Lyman edge, although a complete model based on first principles has yet to be constructed.

While all this is good news, it must be remembered that the Lyman edge problem is only one aspect of the overall theoretical problem. For example, Sincell & Krolik (1997b) [28] have correctly noted that in order to have a standard, viscously heated, optically thick disk be hot enough to drive the local Lyman edge into emission in the innermost radii, the near ultraviolet spectral energy distribution would be close to the long wavelength $F_\nu \propto \nu^{1/3}$ prediction of a standard blackbody disk. This is much harder than observed AGN continuum spectra, assuming one accepts (as I do) that the infrared emission is due to dust and therefore cannot extend underneath the big blue bump. There is still no accretion disk model which satisfactorily fits all the energetically important regions of an AGN spectrum simultaneously.

Variability campaigns in Seyfert galaxies have demonstrated that the OUV and X-ray continua are directly connected through reprocessing. The recent detections of delays between optical and ultraviolet bands in the recent monitoring campaign of NGC 7469 [71] are tremendously exciting, and together with X-rays, may shed light on this problem. Well-sampled, simultaneous, observing campaigns which cover all the energetically important regions of the AGN spectrum (including the
“soft X-ray excess”) are likely to produce significant progress in our understanding of the geometry of the continuum emission regions in the near future.

ACKNOWLEDGMENTS

I have benefited enormously from collaborations and/or conversations with Eric Agol, Robert Antonucci, Robert Goodrich, Ivan Hubeny, Veronika Hubeny, Chia-Ming Hsu, Anuradha Koratkar, Julian Krolik, Paweł Magdziarz, Greg Shields, and Mark Sincell. This work was supported by NSF grant AST 95-29230.

REFERENCES

1. Koratkar, A. 1997, these proceedings
2. Antonucci, R. R. J., Kinney, A. L., & Ford, H. C. 1989, ApJ, 342, 64
3. Koratkar, A. P., Kinney, A. L., & Bohlin, R. C. 1992, ApJ, 400, 435
4. Antonucci, R. 1992, in “Testing the AGN Paradigm”, edited by S. S. Holt, S. G. Neff, & C. M. Urry, 486
5. Antonucci, R., Geller, R., Goodrich, R. W., & Miller, J. S. 1996, ApJ, 472, 502
6. Laor, A. 1990, MNRAS, 246, 369
7. Alloin, D., Pelat, D., Phillips, M., & Whittle, M. 1985, ApJ, 288, 205
8. Cutri, R. M., Wiśniewski, W. Z., Rieke, G. H., & Lebofsky, M. J. 1985, ApJ, 296, 423
9. Courvoisier, T. J.-L., & Clavel, J. 1991, A&A, 248, 389
10. Rauch, K. P., & Blandford, R. D. 1991, ApJ, 381, L39
11. Czerny, B., Jaroszyński, M., & Czerny, M. 1994, MNRAS, 268, 135
12. Korista, K., Ferland, G., & Baldwin, J. 1997, ApJ, 487, 555
13. Kolykhalov, P. I., & Sunyaev, R. A. 1984, Adv. Space Res., 3, 249
14. Kurucz, R. L. 1979, ApJS, 40, 1
15. Tsuji, T. 1976, PASJ, 28, 543
16. Tsuji, T. 1978, A&A, 62, 29
17. Shakura, N. I., & Sunyaev, R. A. 1973, A&A, 24, 337
18. Kriss, G., Krolik, J., Grimes, J., Tsvetanov, Z., Espey, B., Zheng, W., & Davidsen, A. 1997, in “Emission Lines in Active Galaxies: New Methods and Techniques”, edited by B. M. Peterson, F.-Z. Cheng, & A. S. Wilson, 453
19. Mushotzky, R. F., Fabian, A. C., Iwasawa, K., Kunieda, H., Matsuoka, M., Nandra, K., & Tanaka, Y. 1995, MNRAS, 272, L9
20. Nandra, K., George, I. M., Mushotzky, R. F., Turner, T. J., & Yaqoob, T. 1997, ApJ, 477, 602
21. Zheng, W., Kriss, G. A., Davidsen, A. F., Lee, G., Code, A. D., Bjorkman, K. S., Smith, P. S., Weistrop, D., Malkan, M. A., Baganoff, F. K., & Peterson, B. M. 1995, ApJ, 444, 632
22. Zheng, W., Kriss, G. A., Telfer, R. C., Grimes, J. P., & Davidsen, A. F. 1997, ApJ, 475, 469
23. Laor, A., Fiore, F., Elvis, M., Wilkes, B. J., & McDowell, J. C. 1997, ApJ, 477, 93
59. Koratkar, A., Antonucci, R. R. J., Goodrich, R. W., Bushouse, H., & Kinney, A. L. 1995, ApJ, 450, 501
60. Koratkar, A., Antonucci, R., Goodrich, R., & Storrs, A. 1997, ApJ, submitted
61. Blaes, O., & Agol, E. 1996, ApJ, 469, L41
62. Shields, G. A., Wobus, L., & Husfeld, D. 1997, ApJ, submitted
63. Ferland, G. J., & Rees, M. J. 1988, ApJ, 332, 141
64. Ferland, G. J., Korista, K. T., & Peterson, B. M. 1990, ApJ, 363, L21
65. Celotti, A., Fabian, A. C., & Rees, M. J. 1992, MNRAS, 255, 419
66. Barvainis, R. 1993, ApJ, 412, 513
67. Collin-Souffrin, S., Czerny, B., Dumont, A.-M., & Życki, P. 1996, A&A, 314, 393
68. Kuncic, Z., Celotti, A., & Rees, M. J. 1997, MNRAS, 284, 717
69. Nandra, K., & Pounds, K. A. 1994, MNRAS, 268, 405
70. Magdziarz, P., Blaes, O., Zdziarski, A. A., Johnson, W. N., & Smith, D. A. 1997, in “Proceedings of the 4th Compton Symposium”, edited by C. D. Dermer & J. D. Kurfess, in press
71. Wanders, I., et al. 1997, preprint