ALMA imaging of $C_2H$ emission in the disk of NGC 1068

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ABSTRACT

Aims. We study the feedback of star formation and nuclear activity on the chemistry of molecular gas in NGC 1068, a nearby ($D = 14$ Mpc) Seyfert 2 barred galaxy, by analyzing whether the abundances of key molecular species such as ethynyl ($C_2H$), which is a classical tracer of photon dominated regions (PDR), change in the different environments of the disk of the galaxy.

Methods. We used the Atacama Large Millimeter Array (ALMA) to map the emission of the hyperfine multiplet of $C_2H (N = 1–0)$ and its underlying continuum emission in the central $r \approx 35''$ (2.5 kpc) region of the disk of NGC 1068 with a spatial resolution $1'' \times 0''.7$ ($\approx 50–70$ pc).

Results. A sizeable fraction of the total $C_2H$ line emission is detected from the $r \approx 1.3$ kpc starburst (SB) ring, which is a region that concentrates the bulk of the recent massive star formation in the disk traced by the Paα emission complexes imaged by the Hubble Space Telescope (HST). However, the brightest $C_2H$ emission originates from a $r \approx 200$ pc off-centered circunnuclear disk (CND), where evidence of a molecular outflow has been previously observed in other molecular tracers imaged by ALMA. We also detect significant emission that connects the CND with the outer disk in a region that probes the interface between the molecular disk and ionized gas outflow out to $r \approx 400$ pc. We derived the fractional abundances of $C_2H$ assuming local thermodynamic equilibrium (LTE) conditions and a set of excitation temperatures ($T_{ex}$) constrained by the previous multino CO studies of the galaxy. Our estimates range from $X(C_2H) \approx 10^{-6}$ in the SB ring up to $X(C_2H) \approx 10^{-7}$ in the outflow region.

Conclusions. We find that the transient conditions required to fit the high values of $X(C_2H)$ in the outflow are likely due to UV or X-ray irradiated non-dissociative shocks associated with the highly turbulent interface between the outflow and molecular gas in NGC 1068. Although the inferred local timescales are short, the erosion of molecular clouds by the active galactic nucleus (AGN) wind and/or the jet likely resupplies the interface working surface continuously, making a nearly steady state persist in the disk of the galaxy.

Key words. galaxies: individual: NGC 1068 – galaxies: ISM – galaxies: kinematics and dynamics – galaxies: nuclei – galaxies: Seyfert – radio lines: galaxies

1. Introduction

Multiline observations of different molecular tracers can probe the feedback of star formation and/or nuclear activity on the chemistry and energy balance of the interstellar medium of galaxies. There is mounting observational evidence suggesting that the excitation and chemistry of the main molecular species in active galactic nuclei (AGN) are different with respect to those found in purely star-forming galaxies (“starbursts”, SBs) (e.g., Tacconi et al. 1994; Kohno et al. 2001; Usero et al. 2004; Martín et al. 2006; Krips et al. 2008; Aladro et al. 2011, 2013; Imanishi & Nakanishi 2013). Recent progress in the capabilities of millimeter and submillimeter telescopes has made a large set of molecular line transitions from various species, beyond CO line studies, detectable in the nearest sources. These observations can be used to analyze the physical and chemical properties of the different gas components within a galaxy along with its dominant energetics.

The $C_2H$ (ethyl) multiplet at 3 mm has a critical density of $2 \times 10^5$ cm$^{-3}$ at $T_k = 100$ K (not including radiative trapping effects) and hence traces relatively dense gas. However, as for other molecular transitions (e.g., Shirley 2015), the effective excitation density of the $C_2H$ multiplet, which is defined as that producing emission strong enough to be observed (above a $\sim 1$ K km s$^{-1}$ detection threshold in Shirley’s study), can be up to a factor of 10 lower. The emission of $C_2H$ has been detected in a wide variety of molecular cloud environments in our Galaxy from dense Galactic cores ($n(H_2) \geq 10^6$ cm$^{-3}$) to more diffuse clouds ($n(H_2) \geq 10^5$–$10^6$ cm$^{-3}$). In the early studies of Galactic molecular clouds the abundances of ethynyl, defined as $X(C_2H) \equiv N(C_2H)/N(H_2)$, were seen to increase from dense gas, where $X(C_2H) \approx (1–60) \times 10^{-10}$ (Wootten et al. 1980; Huggins et al. 1984; Watt 1983; Watt et al. 1988), to diffuse and translucent gas clouds, where $X(C_2H) \approx 2 \times 10^{-8}$ (Turner et al. 1999). The reason behind this trend is that in diffuse molecular gas, CO can be more easily dissociated by UV photons. This process releases carbon for $C_2H$ formation and increases the abundance of this species in the gas phase, as predicted by stationary-state chemical models (e.g., Sternberg & Dalgarno 1995).

Because of the key role of UV photons, the chemistry of $C_2H$ is generally linked to ionized carbon, so that its abundance is a good tracer of the UV-pervaded outer boundary ($A_V = 1–5$ m$^2$) of...
Galactic molecular clouds, known as photon dominated regions (PDR), where \( X(\text{C}_2\text{H}) \gtrsim 10^{-8} \) (Fuentes et al. 1993; Pety et al. 2005; Beuther et al. 2008; Walsh et al. 2010; Pilleri et al. 2013; Cuadrado et al. 2015). The changes in ethynyl abundances have also been interpreted in terms of chemical evolution in time-dependent models, where chemically younger regions at lower densities would favor higher abundances of \( \text{C}_2\text{H} \) (Li et al. 2012; Pan et al. 2017). Alternatively, Bayet et al. (2011b,a), showed that \( \text{C}_2\text{H} \) can be a tracer of regions influenced by the dissipation of turbulence and waves, heating the gas and accelerating cosmic rays.

The abundance of \( \text{C}_2\text{H} \) has also been found to be comparable to that of molecular species, such as HCN or HCO\(^+\), in a wide variety of external galaxies, i.e., \( X(\text{C}_2\text{H}) \gtrsim 10^{-9}–10^{-8} \). These extragalactic observations include SBs (Henkel et al. 1988; Wang et al. 2004; Meier & Turner 2005; Martín et al. 2006; Aladro et al. 2011; Nakajima et al. 2011; Meier & Turner 2012; Meier et al. 2015), AGNs (Bayet et al. 2011a; Nakajima et al. 2011; Aladro et al. 2013; Martín et al. 2015), low-metallicity objects (Heikkinä et al. 1999), ultraluminous infrared galaxies (ULIRGs; Costagliola et al. 2011) and isolated normal galaxies (Martín et al. 2014). Although the detection of significant \( \text{C}_2\text{H} \) emission in galaxies has been generally linked to PDR chemistry in star-forming regions, other scenarios that point more specifically to outflow-driven chemistry could provide an alternative explanation to the detection of ethynyl in the molecular outflow of Maffei 2 published by Meier & Turner (2012).

NGC 1068 is a prototypical nearby (\( D \approx 14 \) Mpc) Seyfert 2 galaxy, which has been the subject of numerous observational campaigns conducted with single-dish telescopes devoted to studying the fueling of its central region and related feedback using molecular line observations (e.g., Usero et al. 2004; Israel 2009; Kamenetzky et al. 2011; Hailey-Dunsheath et al. 2012; Aladro et al. 2013). NGC 1068 is an archetype of a composite SB/AGN galaxy. High-resolution observations of CO, and of more specific tracers of dense gas such as HCN, HCO\(^+\), CS, and SiO, have proved that interferometers are needed to disentangle the SB/AGN components in the circumnuclear region of this galaxy by spatially resolving the distribution, kinematics, and excitation of molecular gas (Tacconi et al. 1994; Schinnerer et al. 2000; García-Burillo et al. 2010; Krips et al. 2014, 2016; García-Burillo et al. 2014, 2016; Nakajima et al. 2015; Imanishi et al. 2016; Gallimore et al. 2016).

García-Burillo et al. (2014) used the Atacama Large Millimeter Array (ALMA) to map the emission of a set of dense molecular gas tracers (\( \text{CO}(3–2), \text{CO}(6–5), \text{HCN}(4\rightarrow3), \text{CS}(7\rightarrow6) \)) in the central \( r \sim 2 \) kpc of NGC 1068 with spatial resolutions \( \approx 0.3''–0.5'' \) (20–35 pc). The \( \text{CO}(3–2) \) line emission in the ALMA map stems from three regions: 1) the circumnuclear disk (CND); 2) the bar; and 3) the SB ring. Most of the \( \text{CO}(3–2) \) flux in the ALMA map of García-Burillo et al. (2014) comes from the SB ring, which is a two-armed spiral structure that starts from the ends of the stellar bar and forms a molecular gas pseudo-ring at \( r \sim 18'' \) (~1.3 kpc). The SB ring concentrates the bulk of the massive star formation in the galaxy. The CND is an elliptical ring of 350 pc diameter that is off-centered relative to the AGN. In the immediate vicinity of the CND, the \( \text{CO}(3–2) \) emission is detected in an arc-like component on the northeastern side of the disk at distances \( r \sim 5''–6'' \) (400 pc) from the AGN. The anomalous velocities measured both in the arc feature and the CND reveal an AGN-driven massive molecular outflow. García-Burillo et al. (2014) concluded that the molecular outflow is launched when the ionization cone of the narrow line region (NLR) sweeps the disk in the CND and further to the north in the bow-shock arc. More recently, ALMA observations of the CND in the \( \text{CO}(6–5) \), \( \text{HCN}(3–2) \), and \( \text{HCO}^+(3–2) \) lines, carried out with spatial resolutions \( \approx 4–10 \) pc, have been able to isolate and image the dust continuum and molecular line emission from a \( \approx 7–10 \) pc–diameter disk, which represents the submillimeter counterpart of the putative AGN torus of this galaxy (García-Burillo et al. 2016; Imanishi et al. 2016; Gallimore et al. 2016).

Molecular line ratios analyzed by García-Burillo et al. (2014) and Viti et al. (2014) are significantly different in the SB ring and the CND. The change in molecular line ratios inside the CND indicate that radiative and mechanical feedback from the AGN has dramatically changed the physical conditions of molecular gas and chemical abundances of some molecular species in the outflow region (see also Kelly et al. 2017). The SB ring is colder and less dense than the CND and there are also differences in their chemistry (Viti et al. 2014).

The emission of the \( \text{C}_2\text{H} \) multiplet in the disk of NGC 1068 was first observed by Costagliola et al. (2011) and Aladro et al. (2013) with the IRAM 30 m single-dish telescope. In particular, Aladro et al. (2013) used a chemical model to reproduce the average \( \text{C}_2\text{H} \) abundances derived from their data but they were not able to discern whether the \( \text{C}_2\text{H} \) emission was arising from PDRs, dense or shocked gas owing to the insufficient spatial resolution of these single-dish observations (~29'). In this work we use the unique high spatial resolution (~1'') and sensitivity capabilities of ALMA to image the emission of \( \text{C}_2\text{H} \) in the different environments of the NGC 1068 disk. We also use a set of newly developed shock models, gas-phase PDR models, and gas-grain chemical models to account for the abundances derived for \( \text{C}_2\text{H} \).

Hereafter we assume a distance to NGC 1068 of \( D \approx 14 \) Mpc (Bland-Hawthorn et al. 1997); the latter implies a spatial scale of \( \approx 70 \) pc/''.

### 2. Observations

#### 2.1. ALMA data

We observed the emission of the hyperfine multiplet of \( \text{C}_2\text{H} \)\((N = 1) \) in the central \( r \sim 35'' \) (2.5 kpc) region of the disk of NGC 1068 with ALMA during Cycle 2 (project-ID: #2013.1.00055.S; PI: S. García-Burillo). We used Band 3 receivers and a single pointing. We observed one track in August 2015 with 34 antennas with projected baselines ranging from 12 m to 1430 m. The phase tracking center was set to \( \alpha_{2000} = 02^\mathrm{h}42^\mathrm{m}40.771^\mathrm{s}, \delta_{2000} = -00^\circ07'47.94'' \), which is the center of the galaxy according to SIMBAD taken from the Two Micron All Sky Survey – 2MASS survey (Skrutskie et al. 2006). The tracking center is offset by \( \leq 1'' \) relative to the AGN position, \( \alpha_{2000} = 02^\mathrm{h}42^\mathrm{m}40.71^\mathrm{s}, \delta_{2000} = -00^\circ07'47.94'' \) (Gallimore et al. 1996, 2004, 2016; García-Burillo et al. 2014, 2016; Imanishi et al. 2016). The CND and the SB ring of NGC 1068 are both located inside the central \( r \sim 25'' \) (1.8 kpc) of the disk, i.e., well inside the field of view of the 70'' FWHM primary beam of the single-pointed observations.

We used four spectral windows each with a bandwidth of 1.875 GHz and a spectral resolution of \( \approx 0.98 \) MHz (3.4 km s\(^{-1}\)). Two of these windows were placed in the lower side band (LSB) and the other two in the upper sideband (USB). The lowest frequency window of the LSB was centered on the sky frequency 86.3082 GHz. This setup allowed us to observe the whole hyperfine set of lines of \( \text{C}_2\text{H} \). The \( \text{C}_2\text{H} \) (\( N = 1–0 \)) multiplet, hereafter referred to as \( \text{C}_2\text{H} \) (1–0), comprises six hyperfine components grouped in two fine structure groups denoted as...
Table 1. Line components of the C$_2$H($N = 1−0$) multiplet.

| Fine structure group | Hyperfine structure line | Frequency (GHz) | Line strength ($S_{\mu}$) |
|----------------------|--------------------------|----------------|--------------------------|
| Group-I              | $J = 3/2−1/2, F = 1−1$  | 87.284         | 0.17                     |
| Group-I              | $J = 3/2−1/2, F = 2−1$  | 87.317         | 1.67                     |
| Group-I              | $J = 3/2−1/2, F = 1−0$  | 87.329         | 0.83                     |
| Group-II             | $J = 1/2−1/2, F = 1−1$  | 87.402         | 0.83                     |
| Group-II             | $J = 1/2−1/2, F = 0−1$  | 87.407         | 0.33                     |
| Group-II             | $J = 1/2−1/2, F = 1−0$  | 87.447         | 0.17                     |

Notes. Frequencies taken from the Cologne Database for Molecular Spectroscopy (CDMS; Endres et al. 2016).

2.2. Ancillary data

We retrieved the HST NICMOS (NIC3) narrowband (F187N, F190N) data of NGC 1068 from the *Hubble* Legacy Archive (HLA) to obtain calibrated images for the Paα line emission. These images were completely reprocessed with a calibration pipeline as explained in detail in Sect. 2.2 of García-Burillo et al. (2014). The pixel size of the HLA images is 0.72 square.

We also use the CO(3−2) line and continuum maps of NGC 1068 obtained with ALMA during Cycle 0 using Band 7 receivers (project-ID: #2011.0.00083.S; PI: S. García-Burillo) and published by García-Burillo et al. (2014) and Viti et al. (2014).

3. Results

3.1. Continuum map

Figure 1 shows the continuum map derived at 86.3 GHz in NGC 1068. The bulk of the continuum emission stems from a highly structured elongated jet-like feature that emanates from the nucleus and extends out to $r \approx 9''$ along PA $\approx 30^\circ$, the lobe spreading out from the nucleus and extends out to $r \approx 9''$ along PA $\approx 30^\circ$, the lobe spreading out from the nucleus and extends out to $r \approx 9''$ along PA $\approx 30^\circ$.

As our observations do not contain short-spacing correction, we expect to start filtering out an increasing fraction of the line and continuum emissions on scales $>4''$; the latter corresponds to the expected largest angular scale of the emission that can be recovered in the ALMA maps used in this work. This likely affects the faint emission that extends on large scales in the interarm region. As our observations do not contain short-spacing correction, we expect to start filtering out an increasing fraction of the line and continuum emissions on scales $>4''$; the latter corresponds to the expected largest angular scale of the emission that can be recovered in the ALMA maps used in this work. This likely affects the faint emission that extends on large scales in the interarm region. As our observations do not contain short-spacing correction, we expect to start filtering out an increasing fraction of the line and continuum emissions on scales $>4''$; the latter corresponds to the expected largest angular scale of the emission that can be recovered in the ALMA maps used in this work. This likely affects the faint emission that extends on large scales in the interarm region.

3.2. Emission line map

Figure 1 shows the continuum map derived at 86.3 GHz in NGC 1068. The bulk of the continuum emission stems from a highly structured elongated jet-like feature that emanates from the nucleus and extends out to $r \approx 9''$ along PA $\approx 30^\circ$, the lobe spreading out from the nucleus and extends out to $r \approx 9''$ along PA $\approx 30^\circ$, the lobe spreading out from the nucleus and extends out to $r \approx 9''$ along PA $\approx 30^\circ$.

As our observations do not contain short-spacing correction, we expect to start filtering out an increasing fraction of the line and continuum emissions on scales $>4''$; the latter corresponds to the expected largest angular scale of the emission that can be recovered in the ALMA maps used in this work. This likely affects the faint emission that extends on large scales in the interarm region. As our observations do not contain short-spacing correction, we expect to start filtering out an increasing fraction of the line and continuum emissions on scales $>4''$; the latter corresponds to the expected largest angular scale of the emission that can be recovered in the ALMA maps used in this work. This likely affects the faint emission that extends on large scales in the interarm region. As our observations do not contain short-spacing correction, we expect to start filtering out an increasing fraction of the line and continuum emissions on scales $>4''$; the latter corresponds to the expected largest angular scale of the emission that can be recovered in the ALMA maps used in this work. This likely affects the faint emission that extends on large scales in the interarm region. As our observations do not contain short-spacing correction, we expect to start filtering out an increasing fraction of the line and continuum emissions on scales $>4''$; the latter corresponds to the expected largest angular scale of the emission that can be recovered in the ALMA maps used in this work. This likely affects the faint emission that extends on large scales in the interarm region.
showed an arc of gas and dust emission located at the edges of the northeast radio lobe. This gas feature, referred to as the bow-shock arc, has anomalous outward velocities interpreted as the result of the expansion of an AGN-driven molecular outflow in the disk of the galaxy (García-Burillo et al. 2014).

### 3.2. C$_2$H map

Figure 2 shows the C$_2$H(1–0) velocity-integrated intensity map obtained with ALMA in the disk of NGC 1068. We integrated the intensities at each pixel inside the velocity interval defined by $v - v_{\text{sys}} \in [-510, 200]$ km s$^{-1}$. This window encompasses any potential emission stemming from the four strongest hyperfine components of C$_2$H(1–0), which are expected to arise at $v - v_{\text{HEL}} \approx 0$, $-40$, $-290$, and $-310$ km s$^{-1}$. When defining the velocity window we further assumed that the disk rotation of NGC 1068 can cause a maximum velocity shift of the line of $\pm 200$ km s$^{-1}$ around $v_{\text{HEL}}$ (Schinnerer et al. 2000; Krips et al. 2011; García-Burillo et al. 2010, 2014).

Figure 3 shows the C$_2$H(1–0) spectra extracted from a representative set of positions that correspond to strong (>7$\sigma$ in velocity-integrated units) emission knots in the outflow (OUT-I to OUT-III) and the SB region (SB-I to SB-III). Because of the expected non-negligible velocity gradient inside the ALMA beam caused by rotation, non-circular motions, and turbulence, the individual hyperfine components of the fine groups are blended in all the spectra shown in Fig. 3. We nevertheless resolved in the line profiles the emission from the two fine structure groups, which are shifted to each other by $\pm 300$ km s$^{-1}$. Hereafter we define the main and secondary group of lines as the two strongest pairs of hyperfine lines that appear blended around $(v-v_{\text{sys}})_{\text{HEL}} \approx -20$ in group-I, and $-300$ km s$^{-1}$ ($\pm 200$ km s$^{-1}$) in group-II, respectively (see Table 1). No significant emission was detected from the two weakest hyperfine components of C$_2$H in the individual spectra, i.e., outside the range defined by the main and secondary groups. Emission from these satellite lines should lie at $v - v_{\text{sys}} \approx +110$ and $-450$ km s$^{-1}$ ($\pm 200$ km s$^{-1}$) and likely have intensities a factor of $\sim 10$ lower than the strongest line of the C$_2$H multiplet, unless the lines were mostly optically thick, which is a scenario that we discard in Sect. 3.3. Therefore we conclude that the interval defined above to derive the velocity-integrated intensities of the C$_2$H group of lines is not missing any significant emission.

As shown by Fig. 2, a sizeable fraction ($\sim$65%) of the total (spatially integrated) C$_2$H line flux recovered by ALMA stems from the SB ring region. Figure 4 shows the overlay of the Pa$_\alpha$ HST map on the C$_2$H ALMA map. The SB ring concentrates the bulk of the recent massive star formation in the disk traced by the Pa$_\alpha$ emission complexes associated with HII regions. The emission of C$_2$H is mostly prominent at the southwestern end of the SB ring, which is a region hosting an intense star formation activity characterized by strong Pa$_\alpha$ emission. The molecular line emission in the SB ring is very clumpy and it comes from cloud complexes of $\geq 100$ pc-size.

However, the brightest C$_2$H emission originates from the CND. The morphology of the CND displayed by other molecular tracers imaged by ALMA (García-Burillo et al. 2014; Takano et al. 2014; Tosaki et al. 2017) is to a large extent echoed by C$_2$H: the CND appears as an elongated ring with two spatially resolved knots of emission asymmetrically located east and west of the AGN locus (denoted as knots OUT-I and OUT-II, respectively; see Fig. 2). The east and west knots are bridged by fainter emission north of the AGN. We also detected significant emission that connects the CND with the outer disk: in the bow-shock arc region (including knot OUT-III; see Fig. 2) to the northeast, and in two emission lanes east and south of the CND.

Figures 4 and 5 show in color contours the Pa$_\alpha$ and continuum emission, respectively (same as in Fig. 1), on the C$_2$H
Fig. 2. **Left panel:** C$_2$H(1–0) integrated intensity map obtained with ALMA in the disk of NGC 1068. The map is shown in color scale with contour levels $-3\sigma$ (dashed contour), $3\sigma$, $5\sigma$, $7\sigma$, $10\sigma$, $15\sigma$, and $20\sigma$, where $1\sigma = 0.037$ Jy km s$^{-1}$ beam$^{-1}$. We identify the location of the SB ring and the bar. The latter is identified by a representative isophote of the NIR K-band image of 2MASS (Skrutskie et al. 2006). The (red) square markers highlight the positions of the knots SB-I, SB-II, and SB-III along the SB ring. **Right panel:** same as left panel but zooming in on the outflow region, which extends out to $r \approx 6''$ (400 pc), as described by García-Burillo et al. (2014). The outflow region comprises the CND and the bow-shock arc. The position of the AGN ($\alpha_{2000} = 02^h42^m40.71^s$, $\delta_{2000} = -00^\circ00'47.94''$) is highlighted by the star marker. The (red) square markers highlight the positions of the knots OUT-I, OUT-II, and OUT-III. The (red) filled ellipses at the bottom left corners in both panels represent the C$_2$H beam size ($1'' \times 0.7''$ at PA $= 69^\circ$).

Fig. 3. **Upper panels:** C$_2$H emission spectra taken at three representative knots of the SB ring (SB-I, SB-II, and SB-III). **Lower panels:** C$_2$H spectra taken at three representative knots of the outflow region (OUT-I, OUT-II, and OUT-III). The apertures used to extract the spectra correspond to the beam size at the location of each knot ($1'' \times 0.7''$ at PA $= 69^\circ$). The positions of the knots are defined in Fig. 2. The velocity scale for each spectrum has been derived for the strongest C$_2$H line at 87.317 GHz (rest frequency) and is relative to $v_{\text{sys}} = 1136$ km s$^{-1}$. The (red) curves show the composite Gaussian profiles of the best fit to the emission, as described in Sect. 3.3.
ALMA map. Figures 4 and 5 offer evidence that the C$_2$H emission lies at the edges of the AGN ionized nebulousity traced by Paα and radio continuum emission. García-Burillo et al. (2014) analyzed the morphology and kinematics of molecular gas in this region, traced by the CO(3–2) line, and concluded that the NLR gas within the ionization cone is sweeping the ISM of the disk creating a coplanar molecular outflow out to a radius of 400 pc, i.e., a region that comprises the CND and the bow-shock arc feature. The C$_2$H line seems to probe better than CO the interface between the molecular disk and the ionized gas outflow on the same spatial scales. Hereafter we thus refer to the region encompassing the CND and the bow-shock arc as the outflow region.

Figures 6a and b compare the morphology and kinematics, respectively, of molecular gas in the outflow region derived from the C$_2$H(1–0) and CO(3–2) images obtained by ALMA.
The distribution of C$_2$H shows a good correspondence with the brightest emission clumps of CO in the outflow, both in the CND and bow-shock arc, although the emission of CO is comparatively more widespread than that of C$_2$H in this region (see Fig. 6a). Figure 6b compares the emission of C$_2$H and CO in the outflow using the position-velocity ($p-v$) parameter space. With the increase of increasing the signal-to-noise ratio, which is more critical for C$_2$H, the outflow $p-v$ plot diagrams shown in Fig. 6b were derived by averaging the emission along the five axes shown in Fig. 6a, defined by PA = 10$^\circ$, 20$^\circ$, 30$^\circ$, 40$^\circ$, and 50$^\circ$. These axes cover a sizeable fraction of the emission along the molecular outflow, which is oriented along PA = 30$^\circ$ (García-Burillo et al. 2014). Furthermore, with the aim of isolating the non-circular motions associated with the outflow, we subtracted the rotation curve model of García-Burillo et al. (2014) projected along each individual $p-v$ plot before averaging. As illustrated by Fig. 6b, the emission of C$_2$H closely follows the velocities associated with the CO outflow in this region. A significant fraction of the emission of both molecular tracers lies outside the expected range of circular motions attributable to rotation and dispersion: emission is up to 50–200 km s$^{-1}$ redshifted (blueshifted) on the northern (southern) side of the outflow out to $r \approx 6''$ (400 pc). This reflects the sign and the right order of magnitude of the mean-velocity field deviations seen in the CO(3–2) map of García-Burillo et al. (2014).

3.3. C$_2$H line opacities

We performed a fit of the spectra shown in Fig. 3 using a comb of four Gaussian profiles that correspond to the strongest hyperfine components of the C$_2$H(1–0) multiplet. Among the parameters to fit we assumed common widths and excitation temperatures, while optical depth ratios are based on the theoretical line strengths (Tucker et al. 1974). The velocity centroids and optical depths of the strongest component ($\tau_{\text{str}}$) were taken as free parameters. Table 2 lists the results obtained in the fit of the spectra shown in Fig. 3. The fit provides the line parameters for the strongest line ($J = 3/2-1/2, F = 2-1$). We also derived the average secondary-to-main velocity-integrated intensity ratios ($R$) for the CND, the bow-shock arc, and the SB ring. The optical depths derived from the fits of the two spectra located in the CND (OUT-I and OUT-II) are low, i.e., $\tau_{\text{str}} \leq 0.1$ (±0.05). This is in agreement with the average $R$ ratio derived for the CND, $\langle R \rangle = 0.49 \pm 0.05$, which is very close to the optically thin limit for $R \simeq 0.46$. The optical depth derived for the knot OUT-III is nevertheless higher $\tau_{\text{str}} \simeq 1.3$ (±0.6); this result is in line with the higher average $R$ ratio measured in the bow-shock arc, i.e., $\langle R \rangle = 0.70 \pm 0.09$. The optical depths derived for the SB knots (SB-I, SB-II, and SB-III) are low, i.e., $\tau_{\text{str}} \leq 0.1 - 0.2$ (±0.1). The corresponding average $R$ ratio over the SB ring is also low ($\langle R \rangle = 0.50 \pm 0.06$.

We can therefore conclude that the lines of the C$_2$H(1–0) multiplet are mostly optically thin throughout the disk of NGC 1068 with the exception of the bow-shock arc region where opacities may be close to unity$^4$.

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$^4$ In this scenario of low opacities the intensities of the weakest satellite lines would be at most at a level $\leq 2\sigma$ at the position of the strongest emission knot OUT-I.
Table 2. Parameters of C$_2$H line fitting results.

| KNOT   | $a_{2000}$ | $b_{2000}$ | $\delta_{2000}$ | $\eta(T_{\text{ex}} - T_{\text{bg}}) \times \tau_{\text{str}}$ | $FHWM_{\text{str}}$ | $\tau_{\text{str}}$ |
|--------|------------|------------|-----------------|-------------------------------------------------|------------------|-----------------|
| SB-I   | 02h:42m:40.65s | -00'+01'.12s | (2.8 ± 1.0) $\times 10^{-1}$ | 19 ± 8 | 139 ± 22.0 ± 0.6 |
| SB-II  | 02h:42m:40.31s | -00'+01'.09s | (5.3 ± 0.7) $\times 10^{-1}$ | 79 ± 3 | 63 ± 11 0.1 ± 0.2 |
| SB-III | 02h:42m:40.85s | -00'-00'.31s | (6.6 ± 0.5) $\times 10^{-1}$ | -20 ± 2 | 34 ± 10 0.1 ± 0.3 |
| OUT-I  | 02h:42m:40.76s | -00'-00'.48s | (6.8 ± 0.3) $\times 10^{-1}$ | -41 ± 3 | 126 ± 10 0.1 ± 0.1 |
| OUT-II | 02h:42m:40.59s | -00'-00'.49s | (4.1 ± 1.1) $\times 10^{-1}$ | -13 ± 5 | 135 ± 16 0.1 ± 1.9 |
| OUT-III| 02h:42m:40.96s | -00'-00'.45s | (5.4 ± 1.6) $\times 10^{-1}$ | -37 ± 5 | 101 ± 12 1.3 ± 0.7 |

Notes. All the line parameters correspond to the strongest C$_2$H line (J = 3/2 − 1/2, F = 2 − 1) and are derived according to the procedure described in Sect. 3.3 as follows: $T_{\text{ex}}$ is the excitation temperature, $T_{\text{bg}}$ is the background temperature ($\approx 2.73$ K), $\eta$ is the unknown filling factor of the emission, $(v - v_{\text{sys}})_{\text{str}}$ is the velocity centroid of the line relative to $v_{\text{sys}} = 1156$ km s$^{-1}$, $FHWM_{\text{str}}$ is the half width at half maximum of the line, and $\tau_{\text{str}}$ is the estimated peak opacity.

3.4. C$_2$H abundances

In this section we estimate the abundances of C$_2$H relative to H$_2$, defined as $X$(C$_2$H) = N(C$_2$H)/N(H$_2$), both in the SB ring and in the outflow region of NGC 1068.

To estimate the total column densities of C$_2$H, N(C$_2$H), we assumed local thermodynamic equilibrium (LTE) conditions for the sake of simplicity. Under the assumption of LTE conditions, the C$_2$H column densities can be obtained straightforwardly from the intensities of the line. As the velocity-integrated intensities of the C$_2$H multiplet in the optically thick limit are related to the observed intensities via the Boltzmann equation,

$$N = \frac{N_a Z}{g_a e^{-E_a/kT_{\text{ex}}}} \Delta$$

where $N$ is the total column density of the species, $Z$ is the partition function evaluated for $T_{\text{ex}}$, $g_a$ is the statistical weight of the level $u$, and $E_a$ is its energy above the ground state. If the emission is assumed to be optically thin, and if we adopt a filling factor of unity, the column density $N_a$ is related to the observed line intensity,

$$N_a = \frac{8\pi k v^2 I}{hc^2 A_{\text{ul}}},$$

where $I$ is the velocity-integrated line intensity (in K km s$^{-1}$) and $A_{\text{ul}}$ is the Einstein coefficient that corresponds to a particular transition, which in our case corresponds to the $J = 3/2 - 1/2$, $F = 2-1$ line of C$_2$H.

Furthermore, we estimated $N(H_2)$ from the ALMA dust continuum emission map of the galaxy obtained at 349 GHz by García-Burillo et al. (2014) to derive $X$(C$_2$H). These observations well probe the molecular gas column densities for the regions studied here: first, because the dust continuum emission at 349 GHz is mostly optically thin, and, second, because the contribution from atomic hydrogen to the total neutral gas content is negligible in the central $r \approx 30''$ = 2 kpc of the disk (Brinks et al. 1997; García-Burillo et al. 2014). We used a modified blackbody law to derive the dust mass ($M_{\text{dust}}$). In this approach, the fluxes measured in a given aperture, $S_{349\,\text{GHz}}$, can be expressed as $S_{349\,\text{GHz}} = M_{\text{dust}} \times \kappa_{349\,\text{GHz}} \times B_{349\,\text{GHz}}(T_{\text{dust}})/D^2$, where the emissivity of dust at 349 GHz, $\kappa_{349\,\text{GHz}} = 

Table 3. Fractional abundances of C$_2$H in the SB ring and the outflow regions.

|        | SB ring  | Outflow   |
|--------|----------|-----------|
| $T_{\text{ex}}$ = 10 K | (1.3 ± 0.7) $\times 10^{-8}$ | (1.6 ± 0.8) $\times 10^{-8}$ |
| $T_{\text{ex}}$ = 25 K | (9.7 ± 4.0) $\times 10^{-8}$ | (1.1 ± 0.4) $\times 10^{-7}$ |
| $T_{\text{ex}}$ = 50 K | (4.3 $\pm 2.0$) $\times 10^{-7}$ | (4.8 $\pm 2.0$) $\times 10^{-7}$ |
| $T_{\text{ex}}$ = 100 K | (1.8 $\pm 0.8$) $\times 10^{-6}$ | (2.0 $\pm 0.8$) $\times 10^{-6}$ |
| $T_{\text{ex}}$ = 150 K | (4.2 $\pm 2.0$) $\times 10^{-6}$ | (4.7 $\pm 2.0$) $\times 10^{-6}$ |

Notes. Abundances are estimated from the characteristic (1−0) line with gas column densities derived from the dust continuum emission at 349 GHz. We assume LTE conditions and a range of excitation temperatures as explained in Sect. 3.4.

In LTE at temperature $T_{\text{ex}}$ the column density of the upper level $u$ of a particular transition is related to the total column density of the species via the Boltzmann equation,

$$N = \frac{N_a Z}{g_a e^{-E_a/kT_{\text{ex}}}} \Delta$$

where $N$ is the total column density of the species, $Z$ is the partition function evaluated for $T_{\text{ex}}$, $g_a$ is the statistical weight of the level $u$, and $E_a$ is its energy above the ground state. If the emission is assumed to be optically thin, and if we adopt a filling factor of unity, the column density $N_a$ is related to the observed line intensity,

$$N_a = \frac{8\pi k v^2 I}{hc^2 A_{\text{ul}}},$$

where $I$ is the velocity-integrated line intensity (in K km s$^{-1}$) and $A_{\text{ul}}$ is the Einstein coefficient that corresponds to a particular transition, which in our case corresponds to the $J = 3/2 - 1/2$, $F = 2-1$ line of C$_2$H.

In this section the factor is derived from the ratio $=1/(1 + 0.5 + 0.5 + 0.2) = 0.46$. 

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in the SB ring; with a common lower limit of $X(C_2H) = 10^{-6} \times 10^{-8}$, for the lowest temperature we assumed $T_{\text{ex}} = 10$ K. This is close to the formal lower limit on $T_{\text{ex}} \geq 8$ K, which is derived from the values fitted for $\eta_1 \times (T_{\text{ex}} - T_{\text{sp}}) \times \tau_{\text{ex}}$ that are listed in Table 3. This lower limit is formally obtained by assuming a very conservative upper limit on the filling factors in this section. This lower limit is close to the formal lower limit on $X(C_2H)$ in the range $X(C_2H) \geq 8 \times 10^{-7} \times 10^{-6}$. These higher temperatures are specifically required to fit the single-line CO dataset of the outflowing CND (Viti et al. 2014). In particular, Viti et al. (2014) used LTE to conclude that the rotational temperature, $T_{\text{rot}}$, of the CO emitting gas in the CND $\approx 40 - 60$ K. This is also in excellent agreement with the values of the dust temperature required to fit theSED of dust emission in the outflow region out to $r \leq 400$ pc, where $T_{\text{dust}} = 46 \pm 3$ K (García-Burillo et al. 2014). Furthermore Viti et al. (2014) used non-equilibrium radiative transfer schemes (RADEX) to conclude that the kinetic temperature of molecular gas traced by CO lines in the outflow is significantly high, i.e., $\geq 100$ K. In contrast, a similar analysis suggests that kinetic temperatures of molecular gas are comparatively a factor $\geq 2 - 3$ lower in the SB ring.

For temperatures $T_{\text{ex}} \geq 25$ K the derived $C_2H$ abundances lie in the range $X(C_2H) \approx 10^{-5} \times 10^{-6}$. These higher temperatures are specifically required to fit the single-line CO dataset of the outflowing CND (Viti et al. 2014). In particular, Viti et al. (2014) used LTE to conclude that the rotational temperature, $T_{\text{rot}}$, of the CO emitting gas in the CND $\approx 40 - 60$ K. This is also in excellent agreement with the values of the dust temperature required to fit the SED of dust emission in the outflow region out to $r \leq 400$ pc, where $T_{\text{dust}} = 46 \pm 3$ K (García-Burillo et al. 2014). Furthermore Viti et al. (2014) used non-equilibrium radiative transfer schemes (RADEX) to conclude that the kinetic temperature of molecular gas traced by CO lines in the outflow is significantly high, i.e., $\geq 100$ K. In contrast, a similar analysis suggests that kinetic temperatures of molecular gas are comparatively a factor $\geq 2 - 3$ lower in the SB ring.

In order to nail down more accurately how $X(C_2H)$ changes in the different environments of NGC 1068, we would need to estimate $T_{\text{ex}}$ for C$_2$H independently by observing at least 1–2 higher $N$ transitions. However, if we assume that the excitation of C$_2$H follows a pattern similar to the rest of molecular species, the observational evidence discussed above suggests that $T_{\text{ex}}$ is likely a factor of 2–3 higher in the outflow region compared to the SB ring, and, also, that $\geq 10$ K is a conservative lower limit for $T_{\text{ex}}$ in either case. We therefore adopt the following reference values for $X(C_2H)$: $\approx$ a few $10^{-5}$ in the SB ring for $T_{\text{ex}} = 10$ K, and $\approx$ a few $10^{-7}$ in the outflow region for $T_{\text{ex}} = 25 - 50$ K.

4. Origin of the C$_2$H abundances

In this section we aim to account for the range of abundances derived in Sect. 3.4 for C$_2$H in the SB ring and outflow region of NGC 1068 using a set of time-dependent gas-grain chemical models, parametrized shock models, and gas-phase PDR models.

4.1. Chemical models: description

A comprehensive chemical modeling of the C$_2$H molecule is beyond the scope of this observational paper. Moreover, we are aware of the limitations of having observed only one $N$ transition of the C$_2$H multiplet, leading to an uncertain derivation of the fractional abundance. Hence we can only qualitatively discuss the origin of this emission and its potential role as a tracer of distinct gas components. We used and augmented the grid of chemical models from Viti (2017) run with the UCL_CHEM time dependent gas-grain chemical model (Viti et al. 2011), which includes a parametrized shock module (Jiménez-Serra et al. 2008) and a very small grid of PDR models taken from Bell et al. (2006) and Vasta et al. (2010), run with the UCL_PDR code. We assumed for all the models a solar C/O abundance ratio $\sim 0.6$. Below we briefly present the codes, which are more extensively described in the above references.

The code UCL_CHEM is a time dependent gas-grain chemical model that computes the evolution of chemical abundances of the gas and on the ices as a function of time, starting from a diffuse and atomic gas. The code is coupled with the parametrized shock model of Jiménez-Serra et al. (2008). The code is run in two distinct temporal phases denoted as Phase I and Phase II. The initial density in Phase I is $\sim 10$ cm$^{-3}$. The gas is allowed to collapse and reach a high density by means of a free-fall collapse process. In this context the collapse is not meant to represent the formation of protostars, but it is simply a way to compute the chemistry of high density gas in a self-consistent way starting from a diffuse atomic gas, i.e., without assuming the chemical composition at its final density. The temperature during this phase is kept constant at 10 K and the cosmic-ray ionization rates and radiation fields are at their standard Galactic values ($\zeta_0 = 5 \times 10^{-17}$ and 1 Draine field $= 2.74 \times 10^{-16}$ erg cm$^{-2}$ s$^{-1}$, respectively). During Phase I atoms and molecules are allowed to freeze onto the dust grains and react with each other, forming icy mantles. Nonthermal desorption is also included. In Phase II, for the non-shock models, UCL_CHEM computes the chemical evolution of the gas and the icy mantles after either an assumed burst of star formation or an AGN activity episode has occurred. The temperature of the gas increases from 10 K to a value set by the user, and thermal sublimation from the icy mantles occurs. The chemical evolution of the gas is then followed for 10$^7$ yr. In the shock models, Phase II considers a plane-parallel C-shock propagating with a velocity, $V_s$, through the ambient medium.

One of the effects of the AGN activity on the gas is an enhanced cosmic-ray and/or X-ray ionization rate. The grid of models we employed from Viti (2017) uses the cosmic-ray ionization flux to simulate an enhancement in X-ray flux as well. Although this approximation has its limitations in that the X-ray flux ionizes and heats the gas more efficiently than cosmic rays, the chemistry arising from these two fluxes should be similar and the chemical trends are the same for cosmic rays and X-rays for a fixed gas temperature, which is what is assumed in these models (see also Viti et al. 2014).

While some of the models from the UCL_CHEM grid may be considered as representative of PDR models, since the radiation field is enhanced enough to penetrate great depths and hence affect the chemistry, we also considered a very small grid of pure gas-phase PDR models at lower gas density. The latter would be more representative of the diffuse and translucent clouds where C$_2$H has been observed at fractional abundances of $\sim 10^{-5}$ (e.g., Turner et al. 1999) and where the temperature is highly affected by the UV radiation field. These models were run with a self-consistent PDR model: the UCL_PDR code is a 1D time dependent PDR code. It was first developed by Bell et al. (2005, 2006) and further augmented by Bayet et al. (2012) and Priestley et al. (2017). The code assumes a semi-infinite slab geometry and computes the chemistry and temperature as functions of depth and time within the semi-infinite slab. The chemistry and thermal balance are calculated self-consistently at each depth point into the slab and at each time-step, producing chemical abundances, emission line strengths, and gas temperatures as a function of depth and time. The code has been benchmarked (Röllig et al. 2007).

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6 Now publicly available at https://uclchem.github.io and published in Holdship et al. (2017).
computed all our models with solar abundances. UCL_PDR, UCL_CHEM, and shock models we considered. We Tables 4–6 list, respectively, the parameters of the grid of

tables.

4.2. Chemical models: results

Tables 4–6 list, respectively, the parameters of the grid of UCL_PDR, UCL_CHEM, and shock models we considered. We computed all our models with solar abundances. We note that none of the pure gas-phase low density (10^3 cm^{-3}) PDR models (Models 48–51; Table 4), at any time or depth into the cloud, succeed in achieving a fractional abundance of C_2H larger than 10^{-9}. This is consistent with the results of previous PDR models. For example, Turner et al. (1999) succeeded in fitting the C_2H abundance observed in some translucent clouds (10^{-8}) with steady-state models but only by assuming a high non-solar C/O ratio (~0.9). In denser PDRs Cuadrado et al. (2015) found that a C_2H abundance of 10^{-8} can in fact only be matched by a model with a gas density of >10^6 cm^{-3}, radiation field of >10^4 Draine field, and at a depth <2 mag. In other words, only an extremely dense but thin layer of gas can steadily sustain such a high abundance of C_2H. Such a scenario is depicted by some of our models from the UCL_CHEM grid and these models are discussed later.

We now discuss the output of Models 1–44 (Table 5) in the light of the abundances derived for the outflow region. If we assume the temperature of the gas is >50 K then none of the UCL_CHEM models at equilibrium succeed in producing enough C_2H with maximum abundances that reach at most ~3 × 10^{-8}. However, there is a class of models, namely Models 13, 23, 35, where higher abundances, comparable to (or higher than) those observed, can be reached at very early times. These models are all characterized by either a higher (by a factor of 10) than standard UV radiation field, or by a very high cosmic-ray ionization rate (a factor of 5000 higher than standard); the latter would account for the influence of enhanced X-ray irradiation. All models have at least a density of 10^4 cm^{-3}. We plot in Fig. 7 the temporal evolution of C_2H for these models. The increase and decrease of C_2H are due to photodissociation of C_2H_2 and ion-neutral reaction of C_2H with C^+, respectively. While we cannot differentiate between a UV or a X-ray enhanced field, all these models are consistent with those of molecular interfaces between outflows and ambient clouds in high-mass star-forming regions in our Galaxy (Codella et al. 2006; van Kempen et al. 2010; San-José-García et al. 2016), where visual extinctions are low (≤2 mag), gas densities, temperatures, and radiation fields are high (>10^4 cm^{-3}, >100 K, and >10 Draine field, respectively), and the gas is short-lived (~10^{2-3} yr). We find that these transient conditions in the UV or X-ray irradiated gas may be due to non-dissociative shocks associated with the highly turbulent interface between the outflow and molecular clouds. While it is true that Model 23 has a high visual extinction, because of the very high cosmic-ray ionization rate the chemistry of the gas is effectively the same as that of an interface. Although timescales

| Model | \(\zeta_0\) | \(\chi\) | \(n_H\) | \(A_V\) |
|-------|---------|-------|-------|-------|
| 1     | 1.0     | 1.0   | 100.0 | 1e4   |
| 2     | 1.0     | 1.0   | 100.0 | 50    |
| 3     | 1.0     | 1.0   | 100.0 | 5e1   |
| 4     | 1.0     | 1.0   | 100.0 | 5e5   |
| 5     | 1.0     | 1.0   | 100.0 | 5e5   |
| 6     | 1.0     | 1.0   | 100.0 | 5e5   |
| 7     | 1.0     | 1.0   | 100.0 | 5e5   |
| 8     | 1.0     | 1.0   | 100.0 | 5e5   |
| 9     | 1.0     | 1.0   | 100.0 | 5e5   |
| 10    | 1.0     | 1.0   | 100.0 | 5e5   |
| 11    | 1.0     | 1.0   | 100.0 | 5e5   |
| 12    | 1.0     | 1.0   | 100.0 | 5e5   |
| 13    | 1.0     | 1.0   | 100.0 | 5e5   |
| 14    | 1.0     | 1.0   | 100.0 | 5e5   |
| 15    | 1.0     | 1.0   | 100.0 | 5e5   |
| 16    | 1.0     | 1.0   | 100.0 | 5e5   |
| 17    | 1.0     | 1.0   | 100.0 | 5e5   |
| 18    | 1.0     | 1.0   | 100.0 | 5e5   |
| 19    | 1.0     | 1.0   | 100.0 | 5e5   |
| 20    | 500.0   | 1.0   | 100.0 | 5e4   |
| 21    | 500.0   | 1.0   | 100.0 | 5e4   |
| 22    | 5000.0  | 1.0   | 100.0 | 5e4   |
| 23    | 5000.0  | 1.0   | 100.0 | 5e4   |
| 24    | 0.0     | 1.0   | 100.0 | 5e5   |
| 25    | 0.0     | 1.0   | 100.0 | 5e5   |
| 26    | 1.0     | 1.0   | 200.0 | 5e5   |
| 27    | 1.0     | 1.0   | 200.0 | 5e5   |
| 28    | 1.0     | 1.0   | 200.0 | 5e5   |
| 29    | 1.0     | 1.0   | 200.0 | 5e5   |
| 30    | 1.0     | 1.0   | 200.0 | 5e5   |
| 31    | 10.0    | 1.0   | 200.0 | 5e5   |
| 32    | 1.0     | 1.0   | 200.0 | 5e5   |
| 33    | 10.0    | 1.0   | 200.0 | 5e5   |
| 34    | 10.0    | 1.0   | 200.0 | 5e5   |
| 35    | 1.0     | 1.0   | 200.0 | 5e5   |
| 36    | 1.0     | 1.0   | 200.0 | 5e5   |
| 37    | 1.0     | 1.0   | 50.0  | 1e4   |
| 38    | 1.0     | 1.0   | 50.0  | 1e4   |
| 39    | 1.0     | 1.0   | 50.0  | 1e4   |
| 40    | 1.0     | 1.0   | 50.0  | 1e4   |
| 41    | 1.0     | 1.0   | 100.0 | 2e6   |
| 42    | 1.0     | 1.0   | 100.0 | 2e6   |
| 43    | 1.0     | 10^3  | 100.0 | 1e6   |
| 44    | 1.0     | 10^3  | 100.0 | 1e6   |

Notes. Column (1) lists the model number. Columns (2) and (3) list the cosmic-ray ionization rates (\(\zeta\)) and radiation fields (\(\chi\)), respectively, in units of their standard Galactic values (\(\zeta_0 = 5 \times 10^{-17} \text{s}^{-1}\) and 1 Draine field \(= 2.74 \times 10^{-3} \text{erg s}^{-1} \text{cm}^{-2}\)). Column (4) lists the hydrogen gas density (\(n_H\)). For the PDR models the temperature varies with depth and is not listed.

4.3. Shock models: results

The shock models we considered are high non-solar C_2H abundance, high cosmic-ray ionization rate (a factor of 5000 higher than standard); the latter would account for the influence of enhanced X-ray irradiation. All models have at least a density of 10^4 cm^{-3}. We plot in Fig. 8 the temporal evolution of C_2H for these models. The increase and decrease of C_2H are due to photodissociation of C_2H_2 and ion-neutral reaction of C_2H with C^+, respectively. While we cannot differentiate between a UV or a X-ray enhanced field, all these models are consistent with those of molecular interfaces between outflows and ambient clouds in high-mass star-forming regions in our Galaxy (Codella et al. 2006; van Kempen et al. 2010; San-José-García et al. 2016), where visual extinctions are low (≤2 mag), gas densities, temperatures, and radiation fields are high (>10^4 cm^{-3}, >100 K, and >10 Draine field, respectively), and the gas is short-lived (~10^{2-3} yr). We find that these transient conditions in the UV or X-ray irradiated gas may be due to non-dissociative shocks associated with the highly turbulent interface between the outflow and molecular clouds. While it is true that Model 23 has a high visual extinction, because of the very high cosmic-ray ionization rate the chemistry of the gas is effectively the same as that of an interface. Although timescales
Fig. 7. Fractional abundance of C$_2$H for a representative set of chemical models discussed in Sect. 4 that account for the measured abundances of ethynyl in the SB ring and the outflow region of NGC 1068. The C$_2$H abundance is represented as a function of time for gas-grain chemical models (UCL_CHEM) number 13, 23, and 35, and as a function of distance for shock model number 46. This scale corresponds to the distance that the shock travels up to the dissipation length, which is the length scale over which the shock dissipates, beyond which the model is no longer applicable (see Jiménez-Serra et al. 2008). The main parameters of the models, extensively described in Sect. 4 and listed in Tables 4–6, are indicated in each panel. We highlight the range of C$_2$H abundances measured in the SB ring (≈ a few 10$^{-8}$ for $T_{\text{ex}}$ ≃ 10 K) and in the outflow region (≈ a few 10$^{-7}$ for $T_{\text{ex}}$ ≃ 25–50 K).

Table 6. Grid of parametrized shock models.

| Model | ζ (ζ$_0$) | χ (Draine field) | $T_{\text{max}}$ (K) | $n_{\text{H}_{\text{pre-shock}}}$ (cm$^{-3}$) |
|-------|-----------|------------------|----------------------|----------------------------------|
| 45    | 1.0       | 1.0              | 2200                 | 10$^4$                           |
| 46    | 1.0       | 1.0              | 4000                 | 10$^5$                           |
| 47    | 1.0       | 1.0              | 4000                 | 10$^6$                           |

Notes. Column (1) lists the model number. Columns (2) and (3) list the cosmic-ray ionization rates (ζ) and radiation fields (χ), respectively, in units of their standard Galactic values (ζ$_0$ = 5 × 10$^{-17}$ s$^{-1}$ and 1 Draine field = 2.74 × 10$^{-3}$ erg s$^{-1}$ cm$^{-2}$). Columns (4) and (5) list the maximum kinetic temperature ($T_{\text{max}}$) and pre-shock hydrogen density ($n_{\text{H}_{\text{pre-shock}}}$) of the gas, respectively.

Models 1–44, although partly mimicking the effects of nondissociative shocks, did not properly model the passage of the shock. In order to test our hypothesis that the C$_2$H abundances can be reproduced by the erosion of icy mantles due to the passage of nondissociative shocks, we also included in our grid three shock models (45–47; Table 6). In these models we assumed an average shock velocity $V_s$ ≃ 40 km s$^{-1}$. The preshock density and the shock velocity determine the maximum temperature reached (see Table 6). We plot the results of Model 46 in Fig. 7. It is clear that the C$_2$H abundance does indeed rise as the shock penetrates the dissipation length of the gas, and it is close to that found by our LTE analysis if one assumes a 100 K gas, which, especially in the outflow region, is a reasonable gas temperature (García-Burillo et al. 2014; Viti et al. 2014).

With all the necessary caveats in mind, these results suggest that the inferred high abundances of C$_2$H in the outflow could represent the first evidence of a chemical fingerprint of an outflow interface region in an external galaxy. The scenario described above, where UV or X-ray radiation and shocks seem to act in concert to shape the chemistry of the outflow interface may be very short, the continual erosion of dense material by the wind or outflow would resupply the interface, so a nearly steady state may persist in the UV or X-ray illuminated gas across the shock front.
region in NGC 1068, is also similar to that supporting the existence of mechanical dominated regions (MDRs), which is a variant of PDRs in which mechanical action is invoked to explain the spectral energy distribution of CO in some star-forming and active galaxies (Meijerink et al. 2013; Rosenberg et al. 2014, 2015).

We now examine the results of Models 1–44 in the light of the abundances derived for the SB ring. While the derived fractional abundances as a function of temperatures are essentially the same as in the outflow spot, the derived temperatures are comparatively lower in the SB ring (Sect. 3.4) and since C$_2$H may be subthermally excited, the derived fractional abundances may nevertheless correspond to those derived for a gas at a higher (than 10 K) excitation temperature. In this case, the abundances that our models need to match are more of the canonical order of $10^{-8}$. We first look at the models run with a temperature of 50 K (Models 37–40). In fact Model 37, which represents a canonical PDR, with Galactic UV and cosmic-ray ionization fields, a typical molecular cloud gas density of $10^4$ cm$^{-3}$, visual extinction of 1 mag, and temperature of 50 K provides us with a match with an abundance of $10^{-8}$. This may therefore indicate that C$_2$H is simply tracing the gas most exposed to star formation in a region where we know star formation is very active. Other models match this abundance at equilibrium as well: in particular Models 1, 11, 13, and 16, all at $A_V = 1$ mag (hence PDR-like), and differing by a factor of 10 in either radiation field, cosmic-ray ionization field and gas density from Model 37, hence all reasonable PDR models depicting gas affected by active star formation, as expected in a SB ring. However, all these models are run at gas temperatures of 100 K. Moreover, Models 41–44 mimicking the best-fit models from Cuadrado et al. (2015) and, of those, Models 42–44 are able to reproduce the observed abundance of 1–3 x $10^{-8}$, making it difficult to determine whether C$_2$H is tracing a very dense gas subjected to an enhanced radiation field or a more canonical molecular cloud gas. Regardless, the C$_2$H must come from a very low visual extinction gas, which means that the SB ring must be at least partly dominated by PDR gas. Hence for the SB ring, C$_2$H may simply be tracing star formation activity.

Further confirmation of all the favored scenarios discussed above would need a more extensive chemical modeling that explores, for example, a wider range of initial conditions, including initial elemental abundances differing from solar.

5. Summary and conclusions

We used ALMA to map the emission of the hyperfine multiplet of C$_2$H(1–0) and its underlying continuum emission at 86.3 GHz in the central $x = 35''$ (2.5 kpc) region of the disk of the nearby (D = 14 Mpc) Seyfert 2 galaxy NGC 1068 with a spatial resolution 1''0 $\times$ 0''7 ($\simeq$50–70 pc). In an attempt to explain the fractional abundances derived for C$_2$H in the different spatially resolved environments of the galaxy we developed a set of time-dependent chemical models.

We summarize the main results of our study as follows:

- We detected widespread C$_2$H emission in the disk of NGC 1068. Most ($\simeq$65%) of the spatially integrated C$_2$H line flux stems from the $r \simeq 1.3$ kpc SB molecular ring. The SB ring concentrates the bulk of the recent massive star formation in the disk of the galaxy as traced by the strong Pa$_a$ emission of HII regions imaged by HST. The brightest C$_2$H emission originates nevertheless from the $r \simeq 200$ pc outflowing CND. We also detected significant C$_2$H emission bridging the CND with the outer disk in a region that probes the interface between the molecular disk and the ionized gas outflow out to $r \simeq 400$ pc.

- The line ratios measured between the hyperfine components of the C$_2$H multiplet detected throughout the disk of NGC 1068 indicate that their emission is mostly optically thin. We derived the column densities of C$_2$H assuming LTE conditions and a set of excitation temperatures constrained by the previous multimedia CO studies of the galaxy.

- Maps of the dust continuum emission obtained at 349 GHz by ALMA were used to calculate the H$_2$ gas column densities. These were used in combination with the C$_2$H column densities derived from LTE to obtain the fractional abundances of this species in different regions. Our estimates range from $X$(C$_2$H) $\simeq$ a few $10^{-3}$ in the SB ring up to $X$(C$_2$H) $\simeq$ a few $10^{-7}$ in the outflow.

- We performed a preliminary qualitative chemical analysis to determine the origin of the high abundance of C$_2$H in the outflow and the SB ring. In both regions we find that the gas is dense ($\simeq 10^5$ cm$^{-3}$), confirming previous results (García-Burillo et al. 2014; Viti et al. 2014).

- In the outflow we found that fractional abundances of C$_2$H $\simeq$ a few $10^{-7}$ can only be reached at very early times in models where dense molecular gas is heavily irradiated by UV or X-ray photons. In particular, these models are consistent with those of molecular interfaces between outflows and ambient clouds, where visual extinctions are low ($\lesssim$2 mag), gas densities, temperatures, and UV fields or cosmic-ray fields are high ($>10^8$ cm$^{-3}$, >100 K, and >10 Draine field, or >5000 $g_0$, respectively), and the gas is short-lived ($\simeq 10^7$–$10^8$ yr).

- We found that transient conditions in the outflow may be due to UV or X-ray irradiated nondissociative shocks associated with the highly turbulent interface between the outflow and molecular clouds. Although timescales may be very short, the continual erosion of dense material by the wind or jet resupplies the interface, so a nearly steady state persists (e.g., Codella et al. 2006). Shock models seem to confirm that C$_2$H can also be enhanced from icy mantle sputtering.

- In the SB ring, on the other hand, an abundance of C$_2$H $\simeq$ a few $10^{-7}$ is either tracing molecular gas typical of the skin of Galactic giant molecular clouds or is tracing a much denser skin subjected to an enhanced UV field. Both scenarios are consistent with active star formation.

Our chemical modeling is not extensive and has not covered a large enough parameter space of initial conditions to exclude the validity of alternative scenarios. In particular, we cannot exclude that the hydrocarbons chemistry is incomplete in our models or that PAH chemistry (not included) has a significant effect on the gas-phase chemistry of C$_2$H (e.g., Pety et al. 2005); equally a small difference in the C/O ratio in the adopted initial abundances of our models could change our conclusions as this would affect the abundance of all hydrocarbons (Cuadrado et al. 2015). Once further transitions of C$_2$H are obtained, a more comprehensive modeling would be needed to shed light on the origin of the significantly high abundances of C$_2$H in NGC 1068.

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