Super-GZK neutrinos

Andreas Ringwald
Deutsches Elektronen-Synchrotron DESY, Notkestraße 85, D–22607 Hamburg, Germany
E-mail: andreas.ringwald@desy.de

Abstract. In the upcoming decade, observatories for cosmic neutrinos will open up a huge window in energy from $10^7$ GeV to $10^{17}$ GeV, much above the Greisen-Zatsepin-Kuzmin (GZK) cutoff at about $4 \times 10^{10}$ GeV, expected for cosmic protons and nuclei due to inelastic interactions with the cosmic microwave background photons. In this review, we discuss in particular the possibilities to use extremely energetic cosmic neutrinos as a diagnostic of astrophysical processes, as a tool for particle physics beyond the Standard Model, and as a probe of cosmology.

1. Introduction
The future seems bright in extremely high energetic neutrinos (EHEC$\nu$’s). Existing observatories, such as AMANDA [1], ANITA-lite [2], BAIKAL [3], FORTE [4], GLUE [5], and RICE [6] have recently put restrictive upper limits on the neutrino flux in the energy region from $10^7$ to $10^{17}$ GeV (cf. Fig. 1). Furthermore, recent proposals for larger EHEC$\nu$ detectors, such as ANITA [7], EUSO [8], IceCube [9], LOFAR [10], OWL [11], PAO [12], SalSA [13], WSRT [10], together with conservative neutrino flux predictions from astrophysical sources of the observed cosmic rays (CR’s), such as active galactic nuclei, offer credible hope that the collection of a huge event sample above $10^7$ GeV may be realized within this decade (cf. Fig. 1). This will provide not only important information on the astrophysical processes associated with the acceleration of CR’s, but also an opportunity for particle physics beyond the reach of the Large Hadron Collider (LHC). There is even the possibility of a sizeable event sample above $10^{11}$ GeV, with important consequences for cosmology. The corresponding neutrino fluxes may arise from the decomposition of topological defects – relics of phase transitions in the very early universe – into their particle constituents. Moreover, it may be possible to detect the cosmic neutrino background via absorption features in these neutrino spectra. In this review, we will have a closer look at these exciting opportunities.

2. EHEC$\nu$’s as a diagnostic of astrophysical processes
Neutrinos with energies $\lesssim 10^{12}$ GeV propagate essentially without interaction between their source and Earth. Hence, they are a powerful probe of high energy astrophysics, in particular of the conjectured acceleration sites of the CR’s, notably active galactic nuclei (AGN). A paradigm for the acceleration mechanism in the jets of these AGN’s is shock acceleration. Protons and heavier nuclei are confined by magnetic fields and accelerated through repeated scattering by plasma shock fronts. Inelastic collisions of the trapped protons with the ambient plasma produce pions and neutrons, the former decaying into neutrinos and photons, the latter eventually diffusing from the source and decaying into CR protons (cf. Fig. 2 (left)).
Figure 1. Current status and next decade prospects for EHECν physics, expressed in terms of diffuse neutrino fluxes per flavor, $F_{\nu_{\alpha}} + F_{\bar{\nu}_{\alpha}}$, $\alpha = e, \mu, \tau$. Upper limits from AMANDA [1], ANITA-lite [2], FORTE [4], GLUE [5], and RICE [6]. Also shown are projected sensitivities of ANITA [7], EUSO [8], IceCube [9], LOFAR [10], OWL [11], the Pierre Auger Observatory in $\nu_e$, $\nu_{\mu}$ modes and in $\nu_\tau$ mode (bottom swath) [12], SalSA [13], and WSRT [10], corresponding to one event per energy decade and indicated duration. Also shown are predictions from astrophysical CR sources [14], from inelastic interactions of CR’s with the cosmic microwave background (CMB) photons (cosmogenic neutrinos) [14, 15], and from topological defects [16].

A quite conservative estimate of the flux of neutrinos from such astrophysical sources can be made as follows [14]. Assuming that the sources are optically thin, i.e. the neutrons can escape, one may determine the neutron emissivity at the sources from the observed CR spectra [17], taking into account propagation effects, in particular $e^+ e^-$ and pion production through inelastic scattering off the CMB photons. Figure 2 (right) illustrates that both the AGASA and the HiRes data in the $10^{8.6} \div 10^{11}$ GeV range can be fitted nicely under the assumption of a simple power law neutron injection emissivity, $\propto E^{-2.5} (1+z)^3$, of the extragalactic sources, supporting the recent proposal towards a low transition energy, $\sim 10^{8.6}$ GeV, between galactic and extragalactic cosmic rays [22], which is also sustained by chemical composition studies of HiRes data [25]. The neutron injection emissivity is simply related to the neutrino emissivity, and the latter can be translated easily into an expected neutrino flux at Earth. It should be detected very soon, if not already with AMANDA-II, then at least with IceCube (cf. Fig. 1), which therefore can provide significant clues in demarcating the cosmic ray galactic/extragalactic crossover energy [14]. Although the cosmogenic neutrino flux from the inelastic interactions with the CMB photons starts to dominate over the neutrino flux from optically thin cosmic ray sources at energies above a few EeV, it appears to be hard to detect with the EHECν detectors operating in the next decade (cf. Fig. 1).
Figure 2. *Left:* Illustration of shock acceleration in the jet of an active galaxy [14]. *Right:* Best fits [14] to the ultra-high energy cosmic ray spectrum in the energy interval $[10^{8.6}, 10^{11}]$ GeV as observed by Akeno [18]+AGASA [19] and Fly’s Eye [20]+HiRes [21]. The dip from $e^+e^-$ pair production [22, 23] and the bump from Greisen-Zatsepin-Kuzmin [24] (GZK) accumulation are clearly visible in the data and support the simple power law ansatz for the emissivity of the extragalactic sources, in which we have set $z_{\text{min}} = 0.012$, $z_{\text{max}} = 2$, and $E_{i,\text{max}} = 10^{12.5}$ GeV, for the boundaries in redshift and injection energy, respectively. Apparently, this fit undershoots the data for the few highest energy events.

3. EHEC$\nu$’s and physics beyond the Standard Model

Cosmic neutrinos with energies $E_\nu$ above $10^8$ GeV probe neutrino-nucleon scattering at center-of-mass (c.m.) energies above

$$\sqrt{s_{\nu N}} \equiv \sqrt{2 m_N E_\nu} \simeq 14 \text{ TeV } \left( E_\nu / 10^8 \text{ GeV} \right)^{1/2},$$

(1)

beyond the proton-proton c.m. energy $\sqrt{s_{pp}} = 14$ TeV of the LHC, and Bjorken $x \equiv Q^2 / (y s_{\nu N})$ values below

$$x \simeq 2 \times 10^{-4} \left( Q^2 / m_W^2 \right) \left( 0.2 / y \right) \left( 10^8 \text{ GeV} / E_\nu \right),$$

(2)

where $Q^2$ is the momentum transfer squared, $m_W \simeq 80$ GeV the W-boson mass, and $y$ the inelasticity parameter. Under these kinematical conditions, the predictions for $\nu N$ scattering from the perturbative Standard Model (SM) are quite safely under control (cf. Fig. 3 (left)), notably thanks to the input from measurements of deep-inelastic $ep$ scattering at HERA [31, 32]. This makes it possible to search for enhancements in the $\nu N$ cross section due to physics beyond the (perturbative) SM, such as electroweak sphaleron [33] (non-perturbative $B + L$ violation), or Kaluza-Klein, black hole, $p$-brane, or string ball production in TeV scale gravity models [34].

Since the rate of neutrino-initiated showers is proportional to integrated flux times cross section, the non-observation of quasi-horizontal or deeply-penetrating neutrino-induced air showers as reported by, e.g., Fly’s Eye [35], AGASA [36], and RICE [6] can be turned into an upper bound on the neutrino nucleon cross section if a certain prediction for the neutrino flux is exploited [37, 38]. This is exemplified in Fig. 3 (right), which displays the limits on $\sigma_{\nu N}$ from the RICE search results on contained showers [40], for two different assumptions about
the EHECν flux. These bounds are considerably higher than the SM cross section, albeit in the post-LHC energy region. PAO will be able to improve these limits by one order of magnitude [40].

The bounds exploiting searches for deeply-penetrating particles are typically applicable as long as $\sigma_{\nu N} \lesssim 0.5 \div 1$ mb. Models with even higher and more speculative cross sections, $\gtrsim 1 \div 10$ mb, such as electroweak sphaleron production, brane production, or string resonance production, qualify as strongly interacting neutrino scenarios [41, 23], according to which the mysterious EHECR beyond the predicted GZK cutoff [24] at $E_{\text{GZK}} \approx 4 \times 10^{10}$ GeV (cf. Fig. 2 (right)) are initiated by cosmogenic neutrinos. Figure 4 illustrates that a combined fit of the existent data on vertical showers by AGASA and HiRes, as well as of the search results on weakly interacting particles of AGASA and RICE, requires a steep increase within one energy decade around $E_{\text{GZK}}$ by four orders of magnitude [42] – an enhancement which has indeed been proposed within some extensions of the (perturbative) SM.

We have emphasized here the current constraints from EHECν on physics beyond the SM. A more detailed account of the particle physics reach of the planned EHECν observatories can be found elsewhere [47, 48].

4. EHECν’s as a tool to study big bang relics
The existence of topological stable solutions of the field equations (topological defects) is a generic prediction of symmetry breaking (SB) in Grand Unified Theories (GUT’s) and occurs...
Figure 4. The range of the cross section within the 99%, 95% and 90% CL required for a successful strongly interacting neutrino scenario [42]. The lines are theoretical predictions of an enhancement of the neutrino-nucleon cross-section by electroweak sphalerons [43, 44] (short-dashed), p-branes [45] (long-dashed) and string excitations [46] (dotted).
Figure 5. Present (2005) limits on the neutrino flux and projected sensitivity in ten years from now (2015), together with a prediction from topological defects [16] ($m_X = 10^{16}$ GeV, $p = 0$). The absorption dip arising from resonant annihilation of the EHECν’s with big bang relic neutrinos of mass $m_\nu = 0.15$ eV into Z-bosons is clearly visible.

photons from hadronic Z-decay (“Z-bursts”) – which may appear as a CR flux recovery beyond $E_{\text{GZK}}$ and be measured by EUSO, OWL, or LOFAR [16].

5. Conclusions
We are living in exciting times for extremely high energy cosmic neutrinos. There are many observatories under construction, whose combined sensitivity ranges from $10^7$ to $10^{17}$ GeV, the energy scale of Grand Unification. In the likely case that appreciable event samples are collected in this energy range, we can expect a strong impact on astrophysics, particle physics, and cosmology.

References
[1] Ackermann M et al [AMANDA Collab.] 2005 Astropart. Phys. 22 339
[2] Barwick S et al [ANITA Collaboration] 2005 Proceedings of the ARENA Workshop.
[3] Wischnewski R et al [Baikal Collaboration] 2005 Preprint astro-ph/0507698
[4] Lehtinen N, Gorham P, Jacobson A and Roussel-Dupre R 2004 Phys. Rev. D 69 013008
[5] Gorham P, Hebert C, Liewer K, Naudet C, Saltzberg D and Williams D 2004 Phys. Rev. Lett. 93 041101
[6] Kravchenko I 2003 Preprint astro-ph/0306408
[7] Gorham P et al [ANITA Collaboration] 2003 NASA Proposal SMEX03-0004-0019.
[8] Bottai S and Giurgola S [EUSO Collaboration] 2003 Proc. 28th International Cosmic Ray Conference, Tsukuba, Japan, 2003, pp. 1113-1116
[9] Bottai S [EUSO Collaboration] 2003 Proc. Incontro Nazionale di Astrofisica delle Alte Energie, Roma, 2003.
[10] Ahrens J et al [IceCube Collab.] 2003 Nucl. Phys. B (Proc. Suppl.) 118 388
[11] Scholten O, Bacelar J, Braun R, de Bruyn A, Falcke H, Stappers B and Strom R 2005 Preprint astro-ph/0508580
[12] Stecker F, Krizmanic J, Barbier L, Loh E, Mitchell J, Sokolsky P and Streitmatter R 2004 Nucl. Phys. B (Proc. Suppl.) 136C 433; Krizmanic J 2005 private communications.
[13] Bertou X, Billoir O, Deligny O, Lachaud C and Letessier-Selvon A 2002 Astropart. Phys. 17 183
[14] Ahlers M, Anchordoqui L, Goldberg H, Halzen F, Ringwald A and Weiler T 2005 Phys. Rev. D 72 023001
[15] Fodor Z, Katz S, Ringwald A and Tu H 2003 JCAP 0311 015
[16] Ahlers M, Anchordoqui L, Goldberg H, Halzen F, Ringwald A and Weiler T 2005 Preprint DESY 05-165.
[17] Waxman E and Bahcall J 1999 Phys. Rev. D 59 023002
Mannheim K, Protheroe R and Rachen J 2001 Phys. Rev. D 63 023003
[18] Nagano M et al 1992 J. Phys. C 18 423
[19] Takeda M et al [AGASA Collab.] 1998 Phys. Rev. Lett. 81 1163; 2003 Astropart. Phys. 19 447
[20] Bird D et al [Fly’s Eye Collaboration] 1993 Phys. Rev. Lett. 71 3401; 1994 Astrophys. J. 424 491; 1995 Astrophys. J. 441 144
[21] Abu-Zayyad T et al [HiRes Collaboration] 2005 Astropart. Phys. 23 157
[22] Berezinsky V, Gazizov A and Grigorieva S 2002 Preprint hep-ph/0204357; 2005 Phys. Lett. B 612 147
[23] Fodor Z, Katz S, Ringwald A and Tu H, 2003 Phys. Lett. B 561 191
[24] Geisen K 1966 Phys. Rev. Lett. 16 748
Zatsepin G and Kuzmin V 1966 JETP Lett. 4 78
[25] Bergman D [HiRes Collab.] 2004 Nucl. Phys. Proc. Suppl. 136 40
[26] Tu H 2004 Preprint DESY-THESIS-2004-018; unpubl.
[27] Kwiecinski J, Martin A and Stasto A 1999 Phys. Rev. D 59 093002
[28] Gaidashev R, Quigg C, Reno M and Sarcevic I 1998 Phys. Rev. D 58 093009
[29] Glick M, Kretzer S and Reya E 1999 Astropart. Phys. 11 327
[30] Kutak K and Kwiecinski J 2003 Eur. Phys. J. C 29 521
[31] Adloff C et al [H1 Collaboration] 2003 Eur. Phys. J. C 30 1
[32] Coelho C et al [ZEUS Collaboration] 2003 Eur. Phys. J. C 32 1
[33] Aoyama H and Goldberg H 1987 Phys. Lett. B 188 506
[34] Berezinsky V and Kachelriess M 2001 Phys. Rev. D 64 012244
[35] Tkachev I, Khlebnikov S, Kofman L and Linde A 1998 Phys. Lett. B 440 262
[36] Bhattacharjee P and Sigl G 2000 Phys. Rept. 327 109
[37] Bhattacharjee P, Hill C and Schramm D 1992 Phys. Rev. Lett. 69 567
[38] Aloisio R, Berezinsky V and Kachelriess M 2004 Phys. Rev. D 69 094023
[39] Birkel M and Sarkar S, 1998 Astropart. Phys. 9 297
[40] Fodor Z and Katz S 2001 Phys. Rev. Lett. 86 3224
[41] Sarkar M and Toldra R 2002 Nucl. Phys. B 621 495
[42] Barbot C and Drees M 2003 Astropart. Phys. 20 5
[43] Weiler T 1982 Phys. Rev. Lett. 47 234
[44] Singh S and Ma C 2003 Phys. Rev. D 67 023506
[45] Ringwald A and Wong Y 2004 JCAP 0412 005
[46] Eberle B, Ringwald A, Song L and Weiler T 2004 Phys. Rev. D 70 023007
[47] Barenboim G, Mena Requejo O and Quigg C 2005 Phys. Rev. D 71 083002
[48] D’Olivo J, Nellen L, Sahu S and Van Elewyck V 2005 arXiv:astro-ph/0507333
[49] Fargion D, Mele B and Salis A 1999 Astrophys. J. 517 725
[50] Weiler T 1999 Astropart. Phys. 11 303
[51] Fodor Z, Katz S and Ringwald A 2002 Phys. Rev. Lett. 88 171101 2002 JHEP 0206 046
[52] Ringwald A, Weiler T and Wong Y, 2005 Phys. Rev. D 72 043008