This is a repository copy of Systematic analysis of inner crust composition using the extended Thomas-Fermi approximation with pairing correlations.

White Rose Research Online URL for this paper: https://eprints.whiterose.ac.uk/171609/

Version: Accepted Version

Article:
Shelley, Matthew and Pastore, Alessandro orcid.org/0000-0003-3354-6432 (Accepted: 2021) Systematic analysis of inner crust composition using the extended Thomas-Fermi approximation with pairing correlations. Physical Review C - Nuclear Physics. ISSN 1089-490X (In Press)

Reuse
Items deposited in White Rose Research Online are protected by copyright, with all rights reserved unless indicated otherwise. They may be downloaded and/or printed for private study, or other acts as permitted by national copyright laws. The publisher or other rights holders may allow further reproduction and re-use of the full text version. This is indicated by the licence information on the White Rose Research Online record for the item.

Takedown
If you consider content in White Rose Research Online to be in breach of UK law, please notify us by emailing eprints@whiterose.ac.uk including the URL of the record and the reason for the withdrawal request.
Systematic analysis of inner crust composition using the extended Thomas-Fermi approximation with pairing correlations

Matthew Shelley* and Alessandro Pastore†
University of York, Heslington, York, YO10 5DD, United Kingdom,
(Dated: February 25, 2021)

We perform a systematic investigation of the chemical composition of the inner crust of a neutron star, using the extended Thomas-Fermi approximation, the Strutinsky integral correction for shell effects, and the BCS approximation for pairing. 15 Skyrme functionals were selected, which cover the range of values of important bulk properties of infinite nuclear matter, while also having pure neutron matter (PNM) equation of states (EoS) with varying degrees of stiffness. We find that a functional’s low-density PNM EoS is correlated with the number of protons found in the inner crust’s nuclear clusters and, in the lower-density region of the inner crust, with the pressure.

I. INTRODUCTION

Due to a strong pressure gradient, the matter within a Neutron Star (NS) organises itself into different layers [1]. Near the surface of the star lies the crust, where neutrons and protons are arranged in nuclear clusters in a crystalline structure, that transitions gradually into a liquid phase [2] approaching the core of the star [3].

Although the crust only accounts for a small fraction of the mass of the NS, it plays a major role in a variety of astrophysical phenomena, including the r-process [4], short gamma-ray burst precursors caused by resonant shattering [5], soft gamma-ray repeater giant flares [6, 7], and thermal relaxation in soft x-ray transients. The equation of state (EoS) of the crust is also believed to influence many properties of NS, such as the moment of inertia [8] which influences pulsar glitches, transport phenomena within the star [9–11], the relation between the radius and tidal deformability in low-mass NS [12], and the value of the second Love number, $k_2$ [13, 14]. Therefore it is crucial to have a reliable model of the crust of a NS, in order to achieve accurate descriptions of these phenomena.

As discussed in Ref.[1], current models predict that the crust has a crystalline structure, comprising neutron-rich nuclei surrounded by ultra-relativistic electrons [15–17]. The crust can be further separated into the inner and outer parts. The outer crust is the lower-density part of the crust, where neutrons are found only in bound states. The inner crust begins above baryonic densities of $n_b \approx 10^{11} \text{ g cm}^{-3}$, where neutrons begin to drip out of the nuclei, forming a gas.

Since the pioneering work of Negele and Vautherin [18], many theoretical models have been developed to study the properties of the inner crust. The presence of a neutron gas dramatically changes the properties of the clusters [19], in contrast to the outer crust, whose nuclei are isolated. Consequently, determining the composition of the inner crust requires a simultaneous treatment of the clusters and the neutron gas. To simplify this task, it is customary to adopt the Wigner-Seitz (WS) approximation, in which the inner crust is decomposed into independent spheres, named WS cells, centered around each cluster, with a radius $R_{WS}$. Each cell is at β-equilibrium, with a certain number of protons, $Z$, and the same number of electrons (under the condition of charge-neutrality) spread through the cell. Using this approximation, the values of $Z$ (the chemical composition) and $R_{WS}$ can be determined at a given baryonic density, $n_b$, by minimising the total energy per particle, $e$, of the cell [20]. This procedure is valid only in the zero-temperature limit, which is applicable to the case of a non-accreting NS.

Since the density of electrons in the crust is essentially uniform [21], it is possible to calculate their contribution to $e$ analytically [1]. The nuclear contribution is more complex and requires the use of a model. In the literature, several models are used to determine the nuclear binding energy of the system, such as the compressible liquid drop model [22], semi-classical models using the Thomas-Fermi approximation [20, 23–28], or the Hartree-Fock-(Bogoliubov) equations [18, 29].

Ideally, in order to model the EoS of a NS in a unified way, one should use the same model for all of its layers. The ideal choice is a fully microscopic method based on solving the Hartree-Fock-Bogoliubov (HFB) equations [29–31], using an effective nucleon-nucleon interaction [32, 33] adjusted to reproduce selected nuclear observables [34]. A big disadvantage of this approach is how the HFB equations treat continuum states in the inner crust using Dirichlet-Neumann boundary conditions [31, 35, 36], which leads to spurious shell effects. New methods to overcome such a problem have been suggested [37, 38], but no systematic calculations of the WS cells have been performed yet.

To avoid this difficulty, several authors have opted to instead use the extended Thomas-Fermi (ETF) method [24, 26, 39]. Due to its semi-classical nature, the ETF method is not affected by the spurious shell effects of the neutron gas encountered in the standard HFB approach. In Ref. [25], the ETF method was extended to use the Strutinsky integral (SI) correction to recover the important shell effects for the protons in the clus-

---

* mges501@york.ac.uk
† alessandro.pastore@york.ac.uk
This method, named ETFSI, calculates the nuclear energy contribution using parameterised nuclear density profiles, while still using the same energy density functional to generate the fields as in the HFB method. In Ref. [40], the ETFSI method was further developed to take into account the effects of proton pairing correlations.

In a recent series of articles [41, 42], we have performed a systematic comparison between the HFB and ETFSI methods, which have typically yielded different results for the structure of the inner crust. In particular, we observed that by adding to ETFSI the effects of neutron pairing correlations under a simple local density approximation, the resulting energy per particle agrees well with that obtained using the HFB method, apart from the very low density region where the outer-inner crust transition takes place.

Having quantified the agreement of this ETFSI+pairing method with HFB, we can now address an interesting question, namely, why Zirconium isotopes (i.e. clusters with $Z = 40$) are consistently predicted throughout the inner crust [20, 25, 39, 43–45]. To answer this, we perform systematic calculations of the inner crust EoS, selecting several Skyrme functionals whose various infinite nuclear matter (INM) properties [46] cover reasonable ranges. By investigating possible relationships between the INM properties of Skyrme functionals and the proton content and pressure of the WS cells, we aim at providing a better understanding of previous investigations of the structure of the inner crust.

The article is organised as follows: in Sec. II, we outline the main features of the ETFSI+pairing method, and in Sec. III we discuss the Skyrme functionals selected for this work. In Sec. IV, we present results for the chemical composition and EoS of the inner crust, and we give our conclusions in Sec. V.

II. THE ETFSI+PAIRING METHOD

We now briefly review the ETFSI+pairing method, used to study the structure of the inner crust. For a more detailed discussion on the ETFSI method we refer the reader to Refs. [20, 25].

In this article, we consider only cold non-accreting NS, and so we neglect the effects of temperature [47, 48]. Under the WS approximation, for a fixed baryonic density $n_b$ in the inner crust, we minimise the total energy per particle of the WS cell

$$
e = e_{Sky} + e_e - Y_p Q_{n,\beta}.
$$

The last term accounts for the mass difference between neutrons and protons, $Q_{n,\beta} = 0.782$ MeV. $Y_p = Z/A$ is the proton fraction in the cell, where $A$ is the total number of baryons in the cell.

The nuclear part of the energy per particle is expressed as functional of local densities as [33]

$$
e_{Sky} = \frac{1}{A} \int E \left( n_q(r), \tau_q(r), \tilde{J}(r) \right) dr,
$$

where $q = n, p$ stands for neutrons ($n$) and protons ($p$). Considering only time-reversal invariant systems [50], the standard Skyrme functional [51] depends only on a linear combination of the matter densities $n_q(r)$, kinetic densities $\tau_q(r)$, and spin current densities $\tilde{J}(r)$, and on their derivatives.

To calculate $e_{Sky}$, we use the ETFSI+pairing method [25, 40, 42]. In this approach, the densities $n_q(r)$ are not calculated using the wave-functions, as in the HFB method [52], but are instead parameterised using Fermi-Dirac density profiles [25] as

$$n_q(r) = \frac{\rho^{liq}_{q} - \delta_{q,n}\rho^{gas}}{1 + \exp \left( \frac{r - r_{q,n}}{a_q} \right)} + \delta_{q,n}\rho^{gas}.
$$

$\rho^{liq}_{q,n,p}$ are the densities of the neutrons and protons at the centre of the WS cell, $r = 0$, while $\rho^{gas}$ is the neutron density at the cell edge, $r = R_{WS}$. $r_{q,n,p}$ are the radii of the neutrons and protons in the cluster, and $a_{q,n,p}$ are their cluster surface diffusivities. These 7 adjustable parameters are determined by the minimisation of the energy per particle, given in Eq. 1, under the constraints of charge neutrality and $\beta$-equilibrium. We also use the relations presented in Ref. [25], reducing the number of free parameters to 5. In Ref. [26] the authors introduced an exponent for the denominator in Eq. 3, but they concluded that the results are largely unchanged, so we do not consider such a parameterisation.

The ETF method expresses $\tau_q(r)$ and $\tilde{J}(r)$ as a function of the matter densities and their derivatives, up to 4th order in the Wigner-Kirkwood expansion [53]. The full expressions of these quantities are given in Ref. [54].

In this semiclassical approach, shell effects are lost; they are recovered with the Strutinsky integral (SI) correction suggested in Ref. [25]. The main advantages of this ETFSI method over the HFB method are the reliable treatment of the neutron gas [35, 55], and the dramatic reduction in computational cost needed for the energy minimisation.

In the original ETFSI method [20, 25], pairing correlations were not taken into account, despite their fundamental role in describing several important phenomena related to the physics of the inner crust [30, 56–58]. Recent improvements of the model presented in Refs. [40, 42] now allow the inclusion of pairing correlations for both protons and neutrons.
We choose the parameters pairing interaction of the form [59]

\[ v_{q}^{\text{pair}}(r_1, r_2) = v_{0q} \left[ 1 - \eta \left( \frac{n_q(r)}{n_0} \right)^{\alpha} \right] \delta(r_1 - r_2) . \quad (4) \]

We choose the parameters \( \eta = 0.7 \) and \( \alpha = 0.45 \). \( n_0 \) is the saturation density of the functional. We assume that the pairing strength \( v_{0q} \) is the same for neutrons and protons, and we fix it to obtain a maximum pairing gap in pure neutron matter (PNM) of \( \approx 3 \) MeV, as done in Ref. [29]. These choices may appear arbitrary [60], but it was shown in Ref. [42] that a variation of the pairing strength does not impact the resulting chemical composition of the inner crust. To avoid the ultraviolet divergence of the interaction given in Eq.4 [61], we adopt a smooth cut-off in quasi-particle space as defined in Ref. [29].

For functionals of the BSk family [62], we keep the pairing interaction used and developed by the authors.

III. CHOICE OF FUNCTIONALS

According to Ref. [46], more than 240 Skyrme functionals have been published in the literature. Typically, their parameters are adjusted to reproduce the binding energy of some (or all) atomic nuclei, and also some INM properties. There is no standard parameter fitting protocol; various (pseudo)-observables are used in fitting, with widely varying uncertainties. As a consequence, it is difficult to assess the quality of functionals.

The neutron gas in the WS cell is similar to an infinite nuclear system, and comprises the majority of the baryonic matter in the inner crust. We therefore focus on 4 INM properties of the functionals: the saturation density \( n_0 \), the energy per particle \( E_0 \), the symmetry energy \( J \), and its slope \( L \), all evaluated at \( n_0 \). These quantities depend exactly on the EoS of INM [63], and so one can expect that variations in the EoS lead to modifications in the energy per particle of the neutron gas, thus changing the relative energy contributions from baryons and electrons in the WS cell, and ultimately affecting the chemical composition and pressure in the inner crust. In Fig. 1, we show as histograms the distributions of \( n_0, E_0, J \) and \( L \), for all the functionals in Ref. [46].

By inspecting the lower panels of Fig.1, we observe that not all functionals given in Ref. [46] are shown. In our selection process, we imposed two additional constraints on the symmetry energy and its slope, namely \( 30 \leq J \leq 35 \) MeV and \( L = 58 \pm 18 \) MeV. These ranges of values are the ones suggested in Ref. [46]. These constraints are to some extent arbitrary, as there is no consensus on what constitutes a reasonable range of values for \( J \) and \( L \), but they are consistent with the findings of recent analyses [64–68], and with results and associated uncertainties from new chiral EFT calculations [69, 70].

From this group, we finally selected 15 functionals which have a wide range of values for each of \( n_0, E_0, J \) and \( L \): BSk22 and BSk24 [62], KDE [71], KDE0v1 [71], LNS [72], NRAPRii [73] (NRAPR [74] with the spin-orbit strength modified), SS [75], SIV [76], SKRA [77], SKa [78], SLy4 [79], SQMC650 and SQMC700 [80], SkM* [81], and Skz-1 [82]. Their INM properties are labelled in Fig. 1. Five of them — KDE0v1, LNS, NRAPR, SKRA and SQMC700 — are consistent with all the INM constraints presented in Ref. [46], and are named in that reference and hereafter as the CSkP* set.

For our functional selection, we also show the energy

FIG. 1. (Colors online) Each panel represents the distribution of the Skyrme functionals given in Ref. [46] for 4 INM properties. Starting from the left top panel and going clockwise we show \( n_0, E_0, J \) and \( L \). The values of each quantity for the 15 functionals selected in this work are shown as labelled vertical lines. See text for details.
FIG. 2. (Colors online) The EoS of PNM as a function of the neutron density $n_n$ for the 15 functionals selected in this work. The shaded area represents the constraint given in Fig. 4(a) of Ref. [46]. The functionals labelled in the upper left are the Stiff set, and those in the lower right the Soft set. See text for details.

The shaded area represents the constraint given in Fig. 4(a) of Ref. [46]. The functionals labelled in the upper left are the Stiff set, and those in the lower right the Soft set. See text for details.

per particle $e_{PNM}$ in PNM in Fig. 2, for the density range relevant for spherical inner crust calculations. The functionals clearly fall into two distinct families: the set BSk24, KDE, KDE0v1, SII, and SLy4, with a stiff EoS at these low densities (hereafter the Stiff set), and the set BSk22, LNS, NRAPRii, SIV, SKRA, SKa, SQMC650, SQMC700, SkM*, and Skz-1, with a very soft EoS (hereafter the Soft set). On the same figure, we have added a shaded area which corresponds to the range spanned by several ab-initio calculations used to derive the EoS in PNM, and which is discussed in Fig. 4(a) of Ref. [46].

The Stiff set, including SLy4 and BSk24 which have been widely used for NS calculations, are in disagreement with the ab-initio calculations at these densities. In contrast, the majority of the Soft set are in reasonable agreement, apart from Skz-1 which is very soft. However, the ab-initio calculations selected in Ref. [46] do not represent the entire range available in the literature. As shown in Ref.[62], BSk24 and SLy4 are compatible with the error bars provided by chiral effective field theory calculations from Ref [83].

The inner crust does not only comprise neutrons, but also has a significant proton fraction, and so it is interesting to observe the behaviour of the symmetry energy $S$ in the relevant density region and not only at $n_0$. This is shown in Fig. 3, where we again observe a similar grouping into the Stiff and Soft sets as shown in Fig. 2. The reason is quite simple and can be understood from Fig. 1: the vast majority of the functionals have values of $n_0$ and $E_0$ in a very narrow range. This means that all these functionals have a very similar EoS in symmetric nuclear matter (SNM) in the low density region. Since $S$ is just the difference between the EoS in PNM and SNM (within the parabolic approximation [84]), it follows that the pattern observed in PNM repeats here in a very similar way. The exception to this simple rule are SII and SIV. They have extremely low values of $n_0$, shown in Fig. 1, and so a significantly different EoS in SNM compared to the other functionals.

IV. INNER CRUST COMPOSITION

To calculate the EoS and determine the chemical composition of the inner crust, we minimise the total energy per particle $e$ in the WS cell, given in Eq.1, using the method explained in Sec. II and Ref [42]. We cover the range of baryonic densities $n_b \in [0.00025, 0.05]$ fm$^{-3}$, above which non-spherical pasta phases are expected to appear [45], and the range of proton numbers $Z \in [16, 60]$.

We first show results for a few selected baryonic densities $n_b = 0.01, 0.02$ and $0.03$ fm$^{-3}$, for a smaller group of functionals. Figure 4 illustrates the energy per particle as a function of $Z$. We show a re-scaled energy per particle, $e_s$, defined as

$$e_s(n_b) = e(n_b) - e_{PNM}(n_b) - e_{n\text{,pair}}(n_b) \, ,$$

where $e_{n\text{,pair}}$ is the neutron pairing energy per particle. This is purely for visual reasons, so that all values lie within a similar energy range. Since $e_{PNM}$ is independent of $Z$, and $e_{n\text{,pair}}$ is roughly constant with respect to
Z [42], this results in a simple shift in $e$ for a given functional. In the left column are 5 from the Soft set: BSk22, LNS, NRAPRii, SKRA, and SQMC700; in the right column are 4 functionals from the Stiff set: BSk24, KDE0v1, SII, and SLy4, as explained in Sec. III.

In this figure, we clearly see the different behaviour of the Soft and Stiff sets. The Stiff set has a persistent minimum at $Z = 40$, while the Soft set has minima that favour lower $Z$, or that shift towards lower $Z$ as $n_b$ increases, with all favouring $Z = 20$ at $n_b = 0.03$ fm$^{-3}$. The exception in the Soft set is BSk22, which transitions to $Z = 20$ just above $n_b = 0.03$ fm$^{-3}$, and which has one of the stiffest PNM EoS in this set. NRAPRii is unique among all the functionals investigated, in that it favours $Z = 28$ at $n_b = 0.01$ fm$^{-3}$. This is likely related to an issue involving the spin-orbit parameter, $W_0$. See the discussion in Ref. [73] for more details. Below $n_b = 0.01$ fm$^{-3}$, nearer the transition region between the outer and inner crust, other finite-size effects not considered in our ETFSI+pairing method may become significant, changing the results. The HFB method is preferable in this very low density range.

To better understand the origin of the different minima, we plot in Fig. 5 the smooth part of the total energy per particle, $e_{\text{smooth}}$, defined as

$$e_{\text{smooth}} = e - e_{\text{n.pair}} - (e^{SI}_{\text{p}} + e_{\text{p.pair}})$$

where the last term is the sum of the SI correction with BCS pairing for protons, and the proton pairing energy, as explained in Ref. [40]. This results in a smooth parabola-like shape, with a single minimum. We select 2 representative functionals SQMC700 and BSk24, from the Soft and Stiff sets respectively. The parabola shifts significantly in $Z$ going between the two functionals. Furthermore, as $n_b$ increases, we observe that for SQMC700 the minimum moves from $Z = 32$ to $Z = 28$, while for BSk24 it stays at $Z = 46$.

When the SI correction is included, as in Fig. 4, each functional at each density displays local minima at (semi-)magic $Z$ values between 20 and 50. However, Fig. 5 shows how the global minimum in each case is governed more by the stiffness of the PNM EoS. In this work, we strictly consider no temperature effects. Since the various energy minima are quite close to each other, the inclusion of such effects may change this picture. See discussion in Ref. [47] for more details.

A more systematic study of the evolution of the proton
content of the clusters in the inner crust is illustrated in Fig. 6, for all functionals in the Soft and Stiff sets. For clarity, small vertical offsets have been added to each line, and the Soft and Stiff sets in are shown in separate panels. At very low $n_b$ several values of $Z$ are observed, but above $n_b = 0.02$ fm$^{-3}$ only two minima are observed: $Z = 20$ and $Z = 40$. The minimum at $Z = 40$ seems to be favoured by most functionals at around $n_b = 0.01$ fm$^{-3}$, but at $n_b = 0.03$ fm$^{-3}$ the majority of the Soft set have transitioned to the $Z = 20$ configuration. The only exception to this rough classification is SQMC650, whose $e_{PNM}$ is the highest out of the Soft set, as seen in Fig. 2, and it maintains a $Z = 40$ minimum up to quite a high baryonic density of $n_b \approx 0.04$ fm$^{-3}$.

It is interesting to note that $Z = 20$ is favoured by functionals whose PNM EoS is compatible with ab-initio calculations. This result is also in good agreement with the findings of Ref. [58]. Although further analysis is necessary, our results suggest that a better understanding of the EoS in low-density PNM may help to clarify the chemical composition of the inner crust.

Having calculated the chemical composition of the inner crust, we now study its general properties with the different functionals. In Fig. 7, we show the total energy per particle $e$ (Eq. 1) for the WS cells in the inner crust, obtained through the energy minimisation, as a function
of the baryonic density \( n_b \). The pattern formed by the various Soft and Stiff functionals is almost identical to that seen for \( e_{PNM} \), shown in Fig. 2. This further supports the PNM EoS being the major driver of the inner crust EoS.

Another relevant astrophysical quantity is the proton fraction \( Y_p \) in the inner crust, \( Y_p = Z/A \). It plays a role in the transport properties of a NS and in the determination of the neutrino mean free path in the stellar medium [85]. In Fig. 8, we report \( Y_p \) as a function of \( n_b \) for the same smaller selection of functionals shown in Fig. 4. The same feature already observed in Fig. 6 is clear: the Soft functionals predict WS cells with fewer protons, and have a lower \( Y_p \), while the contrary is true for the Stiff functionals. At around \( n_b = 0.02 \text{ fm}^{-3} \), \( Y_p \) varies by up to a factor of 2 between the two sets. When \( Z \) in a WS cell changes, the baryon content of each WS cell follows this change quite closely. As a consequence, the factor of 2 observed between the two dominant minima in Fig.6 \(( Z = 20, 40) \) is conserved here.

It is worth commenting on the case of BSk24, a functional specifically created to be able to provide a unified description of the NS EoS. This was adjusted under many constraints, including the requirement that it reproduce the LS2 equation of state [86] in PNM, which was calculated using the microscopic Brueckner-Hartree-Fock method. The authors note in Ref. [62] that they focused on supernuclear densities when constraining BSk24 to the LS2 EoS. As a result, the BSk24 PNM EoS is significantly stiffer than LS2 at inner crust densities. Around \( n_b = 0.04 \text{ fm}^{-3} \), LS2 is better approximated by SQMC650, SQMC700, and NRAPRii, than by BSk24.

In Ref. [43, 44], the authors claim that the constraining PNM EoS has little effect on \( Y_p \) in the inner crust, and show that \( J \) is an important INM quantity. However, by inspecting carefully Figs. 2 and 8, we see that BSk24 \(( J = 30 \text{ MeV}) \), SLy4 \(( J = 32 \text{ MeV}) \), and KDE0v1 \(( J = 34.6 \text{ MeV}) \) have a \( Y_p \) largely in the range 0.03 – 0.04. The functionals with a softer PNM EoS — BSk22, LNS, NRAPRii, SKRA, and SQMC700 — have a \( Y_p \) mostly in the range 0.02 – 0.025, despite having \( J \) ranging from 31.3 MeV to 33.4 MeV. We therefore conclude that \( J \) is not the driving factor leading to large variations of \( Y_p \), but it is the equation of state of PNM at low density, or equivalently the symmetry energy \( S \) at low density as shown in Fig.3 and discussed in Sec. III.

The softest functional investigated in Ref. [43, 44], BSk22, which more closely follows LS2 at inner crust densities, gives results favouring an energy minimum at \( Z = 20 \) at intermediate inner crust densities, and lower values for \( Y_p \), \( c \), and \( P \). This means that the results presented in this article are consistent with those in Ref. [43, 44], but given the small variations in the properties of the functionals used, the authors were not able to observe such an interesting correlation.

It is also interesting to note from Fig. 8 that almost all of the CSkP* set, selected in Ref. [46] according to several criteria based on INM properties, tend to favour a very small \( Y_p \) within the crust. The only exception is KDE0v1, from the Stiff set.

In Fig.9, we illustrate the evolution of the radius of the Wigner-Seitz cell, \( R_{WS} \), as a function of the baryonic density, again for the smaller selection of functionals. While in the low density region, \( n_b \in [0.00025, 0.015] \text{ fm}^{-3} \), we observe significant varia-
tions in the sizes of the cells, at higher baryonic densities, we notice that the size of the cell is almost independent of the functional. In particular, we clearly identify again the significant jumps shown in this figure correspond to abrupt changes in the chemical composition of the crust, as seen in Fig. 6.

Finally, in Fig. 10, we illustrate the evolution of the pressure $P$ as a function of the baryonic density $n_b$. 4 functionals from the Stiff set are shown, labelled in the upper left, and 5 from the Soft set, labelled just below on the left. See text for details.

that the pressure depends on the derivatives of the EoS, rather than its value.

V. CONCLUSIONS

In this article we have performed a systematic analysis of the pressure $P$ of the inner crust of a non-accreting neutron star (NS), using the ETFSI+pairing method. By selecting 15 different functionals with different infinite nuclear matter properties, we have illustrated a possible correlation between the a functional's pure neutron matter (PNM) equation of state (EoS) at low density, and the proton content and pressure of the Wigner-Seitz cells. In particular, we have shown that functionals with a soft PNM EoS at low densities tend to favor the $Z = 20$ minimum, as well as a lower proton fraction, and a lower pressure up to around $n_b = 0.025$ fm$^{-3}$, while functionals with a stiff PNM EoS show the opposite behaviour, and favour instead the $Z = 40$ minimum. This result clearly shows the importance of constraining the PNM EoS more tightly at subnuclear densities in the adjustment protocol of Skyrme functionals used to study NS. This is not the standard procedure, since authors tend to focus on the high density trend of the EoS; indeed, this is necessary to prevent the collapse of the NS, and to reproduce its global properties such as the maximum mass and radius [87]. More stringent constraints on functionals obtained by more recent ab-initio calculations at low density should pave the way toward a more reliable unified EoS for neutron star matter.

ACKNOWLEDGMENTS

We thank M. Urban for the interesting discussions that motivated this article. This work was supported by STFC Grant No. ST/P003885/1.

[1] N. Chamel and P. Haensel, Physics of neutron star crusts, Living Reviews in relativity 11, 10 (2008).
[2] A. Mantziris, A. Pastore, I. Vidana, D. Watts, M. Bashkanov, and A. Romero, Neutron star matter equation of state including $d^*$-hexaquark degrees of freedom, Astronomy & Astrophysics 638, A40 (2020).
[3] J. Xu, L.-W. Chen, B.-A. Li, and H.-R. Ma, Locating the inner edge of the neutron star crust using terrestrial nuclear laboratory data, Physical Review C 79, 035802 (2009).
[4] J. M. Lattimer, F. Mackie, D. G. Ravenhall, and D. N. Schramm, The decompression of cold neutron star matter, The Astrophysical Journal 213, 225 (1977).
[5] D. Tsang, J. S. Read, T. Hinderer, A. L. Piro, and R. Bondarescu, Resonant Shattering of Neutron Star Crusts, Physical Review Letters 108, 011102 (2012).
[6] C. Thompson and R. C. Duncan, The soft gamma repeaters as very strongly magnetized neutron stars-i. radiative mechanism for outbursts, Monthly Notices of the Royal Astronomical Society 275, 255 (1995).
[7] T. E. Strohmayer and A. L. Watts, The 2004 hyperflare from sgr 1806–20: Further evidence for global torsional vibrations, The Astrophysical Journal 653, 593 (2006).
[8] J. Piekarewicz, F. J. Fattoyev, and C. J. Horowitz, Pulsar glitches: The crust may be enough, Physical Review C 90, 015803 (2014).
[9] B. Link, R. I. Epstein, and J. M. Lattimer, Pulsar constraints on neutron star structure and equation of state, Physical Review Letters 83, 3362 (1999).
[10] A. Burrows, S. Reddy, and T. A. Thompson, Neutrino opacities in nuclear matter, Nuclear Physics A 777, 356 (2006).
[11] E. F. Brown, Nuclear heating and melted layers in the inner crust of an accreting neutron star, The Astrophysical Journal 531, 988 (2000).
[81] H. Krivine, J. Treiner, and O. Bohigas, Derivation of a fluid-dynamical lagrangian and electric giant resonances, Nuclear Physics A 336, 155 (1980).
[82] J. Margueron, J. Navarro, and N. Van Giai, Instabilities of infinite matter with effective skyrme-type interactions, Physical Review C 66, 014303 (2002).
[83] I. Tews, T. Krüger, K. Hebeler, and A. Schwenk, Neutron matter at next-to-next-to-next-to-leading order in chiral effective field theory, Phys. Rev. Lett. 110, 032504 (2013).
[84] D. Davesne, A. Pastore, and J. Navarro, Extended skyrme equation of state in asymmetric nuclear matter, Astronomy & Astrophysics 585, A83 (2016).
[85] N. Iwamoto and C. Pethick, Effects of nucleon-nucleon interactions on scattering of neutrinos in neutron matter, Physical Review D 25, 313 (1982).
[86] Z. H. Li and H.-J. Schulze, Neutron star structure with modern nucleonic three-body forces, Physical Review C 78, 028801 (2008).
[87] J. Zdunik, M. Fortin, and P. Haensel, Neutron star properties and the equation of state for the core, Astronomy & Astrophysics 599, A119 (2017).