CHARACTERIZING THE YOUNGEST HERSCHEL-DETECTED PROTOSTARS. I. ENVELOPE STRUCTURE REVEALED BY CARMA DUST CONTINUUM OBSERVATIONS

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ABSTRACT

We present Combined Array for Research in Millimeter-wave Astronomy 2.9 mm dust continuum emission observations of a sample of 14 Herschel-detected Class 0 protostars in the Orion A and B molecular clouds, drawn from the PACS Bright Red Sources (PBRS) sample. These objects are characterized by very red 24–70 μm colors and prominent submillimeter emission, suggesting that they are very young Class 0 protostars embedded in dense envelopes. We detect all of the PBRS in 2.9 mm continuum emission and emission from four protostars and one starless core in the fields toward the PBRS; we also report one new PBRS source. The ratio of 2.9 mm luminosity to bolometric luminosity is higher by a factor of ~5 on average, compared to other well-studied protostars in the Perseus and Ophiuchus clouds. The 2.9 mm visibility amplitudes for 6 of the 14 PBRS are very flat as a function of uv distance, with more than 50% of the source emission arising from radii <1500 AU. These flat visibility amplitudes are most consistent with spherically symmetric envelope density profiles with ρ ∝ R−2.5. Alternatively, there could be a massive unresolved structure like a disk or a high-density inner envelope departing from a smooth power law. The large amount of mass on scales <1500 AU (implying high average central densities) leads us to suggest that the PBRS with flat visibility amplitude profiles are the youngest PBRS and may be undergoing a brief phase of high mass infall/accretion and are possibly among the youngest Class 0 protostars. The PBRS with more rapidly declining visibility amplitudes still have large envelope masses, but could be slightly more evolved.

Key words: stars: formation – stars: protostars – techniques: interferometric

1. INTRODUCTION

Stars form as a result of the gravitational collapse of clouds of gas and dust. This process can take place in isolated Bok globules (e.g., B335; Keene et al. 1983; Stutz et al. 2008, 2010) or within fragmented giant molecular clouds (e.g., Orion; Johnstone & Bally 1999; Motte & André 2001; Polychroni et al. 2013). The earliest recognizable phase of the star formation process is the Class 0 phase (Andre et al. 1993), the beginning of which is marked by the formation of a hydrostatically supported protostar within an infalling cloud of gas and dust. The robust identification of the youngest sources is imperative for characterizing the initial conditions at the time of protostar formation, before feedback from the formation process significantly alters local physical conditions (e.g., Arce & Sargent 2006; Offner & Arce 2014) and probe the earliest stages of the collapse of the gas onto nascent protostars (e.g., Foster & Chevalier 1993). The density profile, overall mass, and angular momentum of the initially collapsing envelope will determine the potential for fragmentation, how quickly the protostar may accumulate mass, and the growth of the circumstellar disk.

At the start of protostellar collapse, just prior to protostar formation, there is a theoretical prediction of a short-lived (~1000 yr) first hydrostatic core (FHSC), a phase just before or at the start of the Class 0 phase (e.g., Larson 1969; Commerçon et al. 2012). Several candidate FHSCs have been identified (Enoch et al. 2010; Chen et al. 2010; Pineda et al. 2011; Schnee et al. 2012), but their identification as true FHSCs remains uncertain, given that these objects could simply be very low luminosity protostars (e.g., VeLLOs; Bourke et al. 2006). Since the FHSC phase is thought to be short, it is unlikely that many of the candidates in the nearest star forming regions (i.e., Perseus, Taurus, Ophiuchus) are true FHSCs due to their relatively small populations of protostars. Above all else, it is uncertain if there is truly an FHSC phase and if it can be uniquely distinguished from Class 0 protostars. Nonetheless, detecting and characterizing the youngest protostellar sources are key steps toward understanding the star formation process. To capture short-lived phenomena, like the early Class 0 protostars and FHSCs, it is advantageous to look toward more populous regions of star formation.

The Orion molecular clouds are the nearest regions of active star formation. The Spitzer Orion survey by Megeath et al. (2012) found 488 protostellar candidates amongst a total of ~3000 young stellar objects. A subset of 329 protostars from this sample were selected for observations in the far-infrared as part of the Herschel Orion Protostar Survey (HOPS; e.g., Fischer et al. 2010; Stanke et al. 2010; Ali et al. 2010; Manoj...
et al. 2013). Within the fields observed by HOPS, Stutz et al. (2013, hereafter ST13) serendipitously identified 11 protostars with bright 70 μm and 160 μm emission that were not part of the original Spitzer-selected sample. At 24 μm, these sources were either non-detections (8 sources) or so faint that they were flagged as potential extragalactic contamination in the Spitzer surveys (Megeath et al. 2012; Kryukova et al. 2012). Moreover, within the original HOPS sample, seven protostars had [24 μm]=[70 μm] colors (in log (S/F_r) space) redder than 1.65, consistent with the Herschel-detected sources. ST13 refers to the 18 protostars satisfying the extremely red color criterion as the PACS Bright Red Sources (PBRS).

Analysis of the PBRS spectral energy distributions (SEDs) by ST13, which were augmented by APEX 350 μm and 870 μm mapping, found that these PBRS sources have very cold bolometric temperatures (T bol; Myers & Ladd 1993; 29–45 K) and high ratios of submillimeter (submm) luminosity to bolometric luminosity (Lbol) (0.6%–6.1%). Most PBRS are not detected shortward of 24 μm, but some display faint features in the Spitzer IRAC 4.5 μm band, possibly indicative of shocked H_2 emission associated with outflows. Despite their deeply embedded nature, sources emitting at 70 μm must be self-luminous. For example, starless core models show that the emission would otherwise be too faint to detect at 70 μm (Ragan et al. 2012; ST13). It is important to point out that the PBRS sources are not low-luminosity objects like the VeLLOs (e.g., Bourke et al. 2006) since they have Lbol ranging between 0.65 L_⊙ and 30.6 L_⊙, with the median being ~3 L_⊙. These luminosities are large enough such that they are not dominated by external heating (Dunham et al. 2008).

The characteristics of the PBRS indicate that these protostars could be very young Class 0 sources with very dense envelopes (ST13). There is, however, a degeneracy in the interpretation of protostar SEDs between envelope density and inclination due to bipolar cavities being evacuated by the outflows. The envelope properties were also difficult to study with only the APEX submm data available in ST13, due to the low resolution and blending of the envelope emission with extended cloud structure. Furthermore, the lack emission shortward of 10 μm toward most PBRS, made the inclinations impossible to constrain from SED modeling. The only way to derive more detailed envelope properties, independent of inclination, is to observe these sources with an interferometer.

We have obtained observations of a subset of the PBRS sample with the Combined Array for Research in Millimeter-wave Astronomy (CARMA). There are a total of 19 PBRS, 18 of which were presented in ST13 and 1 additional PBRS will be described in this paper. We have observed 14 PBRS, focusing on the new, Herschel-detected subset of PBRS. We focused on this subset because they had the least amount of complementary data and non-detections at most wavelengths shorter than 70 μm. We observed the PBRS in both the dust continuum and spectral line emission to examine both the envelope and outflow properties of these sources; the outflow results will be presented in a future paper. We discuss the observations in Section 2, our results for the dust continuum emission and model comparison are given in Section 3, we discuss the results within the broader context of star formation in Section 4, and summarize our main conclusions in Section 5.

2. CARMA OBSERVATIONS AND DATA REDUCTION

CARMA is a heterogeneous interferometer array comprised of 23 antennas (6–10.4 m, 9–6.1 m, and 8–3.5 m) located in the Inyo mountains of California. Our observations were carried out with the main, 15-element CARMA array using the 10.4 m and 6.1 m antennas in two configurations. We observed a subset of the PBRS identified in ST13 with CARMA in the D configuration during late 2012 and early 2014. We also followed-up a subset of the sources observed in D-configuration with higher-resolution observations in C-configuration in early 2014. The angular resolutions in D and C configurations were ~5′′ and ~2′′ respectively. The central frequency was 107.77 GHz and four spectral windows were configured for 500 MHz bandwidth to observe the dust continuum; two windows were configured for 8 MHz bandwidth to observe para-NH2D (J = 111 → 101) and C18O (J = 1 → 0); and the two remaining windows had 31 MHz bandwidth to observe 13CO (J = 1 → 0) and 12CO (J = 1 → 0). Two or three sources were observed per track, with further details given in Table 1. The C-configuration observations did not observe para-NH2D and another 500 MHz continuum band was allocated. Generally, we only detect 12CO (J = 1 → 0) and the 2.9 mm continuum; there were a few weak detections in the other lines which will not be discussed further. Our root-mean-squared (rms) sensitivity is typically 0.2 Jy beam−1 channel−1 for the CO (J = 1 → 0) and 1 mJy beam−1 for the continuum data. The data were reduced, edited, and imaged using standard procedures within the MIRIAD software package (Sault et al. 1995). The uncertainty in the absolute flux is estimated to be ~20%. We will only present the continuum results in this paper, the CO outflow data will be presented in an upcoming paper.

3. RESULTS

The observations of λ = 2.9 mm continuum emission enable us to probe the properties of the protostellar envelopes, in terms of mass and density profiles. We will discuss the overall flux densities, visibility amplitudes, and comparison of the visibility amplitudes to radiative transfer models. We also report the data observed toward additional sources located within our field of view, but primarily discuss the PBRS in the main text; the PBRS and non-PBRS are denoted in Table 2.

3.1. Integrated 2.9 mm Dust Continuum Emission

We detect all the observed PBRS sources in the 2.9 mm continuum and deconvolved images using natural weighting are shown in Figure 1. Our observations are sensitive to spatial scales between ~1000 AU and ~10,000 AU. On these scales, most sources have some resolved structure, in terms of extended envelope emission and in the case of HOPS 373, there is a binary source separated by ~4″. Two sources (082012 and 061012) also have companions ~20″ (9400 AU) away. Images of the sources made with Robust weighting factor of ~1 (Briggs 1995) did not reveal significant structure on smaller-scales. The flux densities measured from the deconvolved images are presented in Table 2. The 2.9 mm flux densities of the sample exhibit a relatively large amount of heterogeneity given the extremely red colors selection of the sample. The brightest PBRS is 082012 at 155.6 mJy and the faintest is 119019 at 10.2 mJy. Indeed, 12 of 14 PBRS have flux densities ~30 mJy and their values of Lbol also span an order of magnitude. The combined D- and C-configuration images agree with D-configuration-only flux densities within the statistical uncertainties. Note that we also present an additional PBRS, 135003, that did not appear in ST13. This source was left out from the sample due to the 70 μm FWHM being extended more than the cutoff value 7″.8. More details of this source and
its infrared and submm imaging are given in the Appendix; its inclusion raises the number of PBRS in the Orion clouds to 19.

The strength of dust continuum emission from the PBRS sources prompted us to collect $\lambda \sim 3$ mm flux densities from the literature of other Class 0 or Class I protostellar sources observed with interferometers for comparison (Table 3; Looney et al. 2000; Arce & Sargent 2006; Tobin et al. 2011). These observations had comparable resolution and sampling of the $uv$-plane. To match the 2.9 mm flux densities better, we have scaled the flux densities of the comparison sources. We have done the scaling by assuming that the relative flux densities only depend on the dust opacity spectral index ($\beta$) and the function $F_{\lambda} \propto \lambda^{-2(\beta+1)}$. This assumption is reasonable given the similar wavelengths of the samples. With the further assumption that $\beta \sim 1$, the scaling factors for the 2.7 mm flux densities and 3.4 mm
continuum images at 2.9 mm for the observed PBRS sample and sources within the field of view. While many appear mostly round, some have marginally resolved emission and Table 3 shows that the peak flux densities are less than the integrated flux densities. The sources in (a) are a combination of D and C configuration data and the sources with only D-configuration data are shown in (b). The contours are $[-3, 3, 6, 9, 12, 15, 20, 25, 30, 35, 40, 45, 50, 60, ...] \times \sigma$ for all sources except 093005, 090003, and 082012 where the contours are $[-3, 3, 6, 15, 25, 45, 65, 85, ...] \times \sigma$; $\sigma$ is given for each source in Table 2.

The comparison sources span the range of observed $L_{\text{bol}}$ for the PBRS, but most PBRS have lower $L_{\text{bol}}$ values, comparable to those in Tobin et al. (2011). They, however, have $L_{2.9\text{mm}}$ values that are comparable to the Looney et al. (2000) sources, which are among the brightest nearby protostars a millimeter wavelengths (e.g., NGC 1333 IRAS 4A, NGC 1333 IRAS2A, IRAS 16293-2422) and are more luminous than most PBRS.

Furthermore, the PBRS have among the largest values of $L_{2.9\text{mm}}$ and $L_{2.9\text{mm}}/L_{\text{bol}}$ ratios. This behavior is true at all luminosities, but especially evident at $L_{\text{bol}} \sim 1 L_\odot$. The non-PBRS sources in our observations generally have higher $L_{\text{bol}}$, higher $T_{\text{bol}}$, and lower $L_{2.9\text{mm}}/L_{\text{bol}}$ ratios; however, the source HOPS 68 does intermingle with the PBRS in Figure 2. Note that the results do not significantly change whether or not scaling is applied to the literature data.

Also evident in Figure 2 is the lack of a clear relationship between $L_{\text{bol}}$ and $L_{2.9\text{mm}}$. This indicates that the millimeter emission is decoupled from the central source properties (central source refers to both the protostar and accretion processes generating luminosity). The PBRS are tracing a new region of parameter space with their large amounts of circumstellar material traced by the 2.9 mm flux densities and lower values of $L_{\text{bol}}$.

Figure 1. Continuum images at 2.9 mm for the observed PBRS sample and sources within the field of view. While many appear mostly round, some have marginally resolved emission and Table 3 shows that the peak flux densities are less than the integrated flux densities. The sources in (a) are a combination of D and C configuration data and the sources with only D-configuration data are shown in (b). The contours are $[-3, 3, 6, 9, 12, 15, 20, 25, 30, 35, 40, 45, 50, 60, ...] \times \sigma$ for all sources except 093005, 090003, and 082012 where the contours are $[-3, 3, 6, 15, 25, 45, 65, 85, ...] \times \sigma$; $\sigma$ is given for each source in Table 2.

flux densities are 0.8 and 1.6 respectively. We have converted all the flux densities to 2.9 mm luminosities ($L_{2.9\text{mm}}$) using the distances provided in Table 3 and assuming a bandwidth of 0.11 mm (4 GHz). We plot $L_{2.9\text{mm}}$ versus $L_{\text{bol}}$ and the ratio of $L_{2.9\text{mm}}$ to $L_{\text{bol}}$ versus $L_{\text{bol}}$ for all the data in Figure 2.

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3.2. Circumstellar Masses

The integrated flux densities of the protostars enable us to directly probe the circumstellar mass associated with the protostars, without significant contributions from the surrounding molecular cloud. In this case, the interferometer filtering works to our advantage by separating the envelope emission from the surrounding background cloud. To convert a flux density into a mass, we assume that the emission is optically thin and isothermal, and apply the equation

$$M = \frac{D^2 F_\lambda}{\kappa_{2.9\,\text{mm}} B_\lambda(T_{\text{dust}})},$$

where $B_\lambda$ is the Planck function. We have assumed that $T_{\text{dust}} = 20$ K, $\kappa_{2.9\,\text{mm}} = 0.00215$ cm$^2$ g$^{-1}$ using Ossenkopf & Henning (1994, Table 1, Column 5) extrapolated to 2.9 mm, and $D = 420$ pc. The extrapolation to 2.9 mm uses the dust opacity spectral index ($\beta$) of the Ossenkopf & Henning (1994) dust model between 700 $\mu$m to 1.3 mm which has $\beta = 1.78$. The opacity given is the dust+gas opacity, assuming a gas-to-dust ratio of 100.

The calculated masses are given in Table 2, the uncertainties given are statistical only (not including the uncertainty in absolute flux calibration) and the masses themselves are likely only valid at the order of magnitude level given the assumptions. There may be optically thick regions of the envelope, but those are on scales of order a few hundred AU and will make only a small contribution to the overall mass.

With these assumptions, all the PBRS sources (except 119019) have more than 1 $M_\odot$ of surrounding material, with the largest being almost 10 $M_\odot$. These masses are reflected in the high integrated flux densities observed toward these protostars. We also note that the masses are systematically larger than those calculated by ST13. The masses calculated in ST13, however, are from modified blackbody fits to the emission from 70 $\mu$m to 870 $\mu$m and the ST13 modified blackbody fits systematically underpredict the 870 $\mu$m flux densities. The underprediction of the 870 $\mu$m flux densities likely results from fitting a single temperature to data that reflect a superposition of temperatures and span an order of magnitude in wavelength. If masses were calculated directly using the 870 $\mu$m flux, closer agreement is expected. Several non-PBRS have masses listed in Table 2 that are comparable to the mass of the PBRS. Many of these sources, however, have higher $L_{\text{bol}}$ and $T_{\text{bol}}$, suggesting that the dust temperatures could be larger and by extension the masses are overestimated. We assumed $T_{\text{dust}} = 20$ K, and the actual masses will be different by the ratio $T_{\text{dust}}/20$ K.

The estimated masses will also change if we assume a different dust opacity law; Ossenkopf & Henning (1994) has $\beta = 1.78$ at millimeter wavelengths (Table 1, Column 5). If we instead assume $\beta = 1$ and use the normalization of $0.1(\nu/1200\,\text{GHz})^\beta$ from Beckwith et al. (1990), the masses would be a factor of 4.5 lower.

3.3. Visibility Amplitudes of 2.9 mm Continuum

The integrated flux densities are only one aspect of the continuum data, the visibility amplitudes as a function of $uv$ distance/baseline length can reveal more about the source structure than the deconvolved images alone. We show the visibility amplitudes for all detected sources in Figure 3. How slowly (or rapidly) the amplitude decreases with increasing $uv$ distance
reveals how concentrated the emission is toward a particular source, in addition to structural changes in the emitting material. Similar to the order of magnitude span in 2.9 mm flux density, the visibility amplitudes profiles themselves are quite varied but generally fall within two groups. About half of the observed PBRS have amplitudes that drop quickly with increasing uv distance (HOPS 373, 082012, 302002, 119019, HOPS 372, 019003A, 061012), meaning that there is more emission on larger spatial scales relative to small spatial scales. The other half of the sample have amplitudes that are flat or slowly decreasing larger than those from Looney et al. (2000). Note that we have rescaled the flux densities from the literature to account for the spectral slopes, assuming $\beta = 1,$ the 3.4 mm flux densities are increased by a factor of 1.6 and the 2.7 mm flux densities are decreased by a factor of 0.8. Note that the scaling does not affect the result of the PBRS having among the highest $L_{bol}/L_{bol}$. The visibility amplitude data. The goal is to determine what models to obtain qualitative results for the interpretation of the visibility amplitude data. The goal is to determine what

| Source          | $S_v$ (mJy) | Distance (pc) | $L_{bol}$ ($L_\odot$) | $\lambda$ (mm) |
|-----------------|-------------|---------------|------------------------|----------------|
| L1448 IRS3B     | 134.6 ± 3.9 | 230           | 5.0                    | 2.7            |
| L1448 IRS3C     | 31.7 ± 4.1  | 230           | 1.0                    | 2.7            |
| NGC 1333 IRAS2A | 82.8 ± 4.0  | 230           | 20.0                   | 2.7            |
| NGC 1333 IRAS4A | 544.2 ± 13.6| 230           | 5.8                    | 2.7            |
| NGC 1333 IRAS4B1| 180.3 ± 7.9 | 230           | 3.8                    | 2.7            |
| NGC 1333 IRAS4B2| 49.8 ± 5.5  | 230           | 3.8                    | 2.7            |
| L1551 IRS5      | 173.3 ± 7.5 | 140           | 24.5                   | 2.7            |
| VLA 1623        | 72.1 ± 6.8  | 125           | 1.0                    | 2.7            |
| IRAS 16293-2422 | 1017.9 ± 26.5| 125          | 15.0                   | 2.7            |
| Perseus 5       | 4.0 ± 1.0   | 230           | 0.5                    | 3.4            |
| HH211           | 43.4 ± 2.0  | 230           | 1.2                    | 3.4            |
| L1527           | 32.6 ± 6.0  | 140           | 1.7                    | 3.4            |
| HH270VLA1       | 6.3 ± 1.0   | 420           | 7.0                    | 3.4            |
| RNO43           | 9.4 ± 3.5   | 420           | 12.5                   | 3.4            |
| IRAS 16253-2429 | 1.8 ± 1.0   | 125           | 0.25                   | 3.4            |
| HH108MMS        | 10.3 ± 2.3  | 300           | 0.7                    | 3.4            |
| HH108IRS        | 35.7 ± 4.9  | 300           | 8.0                    | 3.4            |
| L1165           | 12.1 ± 1.7  | 300           | 13.0                   | 3.4            |
| L1152           | 15.2 ± 2.8  | 300           | 1.0                    | 3.4            |
| L1157-CARMA     | 55.1 ± 8.3  | 300           | 5.0                    | 3.4            |
| RNO43           | 26.6 ± 6.0  | 420           | 12.5                   | 2.7            |
| HH114MMS        | 110 ± 20    | 420           | 26.0                   | 2.7            |
| IRAS03262+3035  | 61 ± 12    | 230           | 1.2                    | 2.7            |
| IRAS04239+2436  | 14 ± 2     | 140           | 1.1                    | 2.7            |
| IRAS20582+7724  | 45 ± 9     | 200           | 4.0                    | 2.7            |
| IRAS21000+7811  | 62 ± 13    | 200           | 1.0                    | 2.7            |

Note. The collection of 2.7 mm data at the top are from Looney et al. (2000), the 3.4 mm data are from Tobin et al. (2011), and the collection of 2.7 mm data at the bottom are from (Arce & Sargent 2006).
Figure 3. Visibility amplitude vs. uv distance plots for all the sources. Flatter visibility amplitudes with increasing projected uv distance indicate that the flux density is dominated by compact, unresolved emission (e.g., 093005, 090003, 082005, 091015, 091016). Visibility amplitudes that are decreasing rapidly with increased uv distance indicate that there is more flux on larger spatial scales, relative to a compact unresolved component (e.g., 082012, 302002, HOPS 373, HOPS 223). The light dashed line in each plot is the expected visibility amplitude that would be measured from noise alone. Sources only observed in D-configuration with shorter uv distances are shown in (b).

density profiles are consistent with the data and if a compact, unresolved source is a necessary component for the models to fit the data. We use the Hyperion code (Robitaille 2011) to perform the radiative equilibrium calculations and produce ray-traced images of 2.9 mm continuum emission, with a $5.0 M_\odot$ envelope, $1 L_\odot$ central protostar, 10,000 AU outer radius, and radial density profiles of $\rho \propto R^{-1.5}$, $R^{-2.0}$, and $R^{-2.5}$, and a 50 AU radius embedded disk with $M_{\text{disk}} = 0.0 M_\odot$, 0.01 $M_\odot$, and 0.1 $M_\odot$. We also ran envelope models using the density structure for a rotationally flattened, infalling envelope (CMU envelope; Ulrich 1976; Cassen & Moosman 1981). For the CMU models, we explored the same disk masses, but we used four centrifugal radii ($R_C = 50$ AU, 100 AU, 300 AU, and 500 AU) and assumed that the disk radius was equal to $R_C$. $R_C$ is the radius at which infalling material can be rotationally supported due to conservation of angular momentum. The overall envelope masses of the CMU models were the same as those of the power-law envelopes. The inclination of the system only has a minor effect on the visibility amplitudes and, we assume an inclination angle of 60° for simplicity.

We use the dust opacities calculated by Ormel et al. (2011) for icy silicate grains and bare graphite grains grown for a period of $3 \times 10^5$ yr. These dust opacities are similar to those of Ossenkopf & Henning (1994, Table 1, Column 5), but are calculated down to $\lambda \sim 0.1 \mu m$ and include scattering properties. The dust opacity spectral index ($\beta$) of the Ormel et al. (2011) models, however, is $\sim 2$ at submm and millimeter wavelengths, consistent with interstellar-medium-sized dust grains. This steep $\beta$, however, results in a very low dust opacity at 2.9 mm and very faint envelope emission. Therefore, we have altered the dust opacity model and at wavelengths greater than 90 $\mu m$ we transition to the Ossenkopf & Henning (1994, Table 1, Column 5) dust opacity model. This model has $\beta \sim 1.78$, yielding $\kappa_{2.9\,\mu m} = 0.00215$, producing 2.9 mm millimeter fluxes more consistent with our observations.

We could have simply increased the envelope masses such that the flux densities were consistent with our data. The Ormel et al. (2011) dust opacities, however, are a factor of 2.35 lower than Ossenkopf & Henning (1994, Table 1, Column 5). Thus, it would have been necessary to increase envelope masses to
\(~12 \, M_\odot\), which may be unrealistically large for many of our sources. Furthermore, the larger masses would increase the envelope opacity at shorter wavelengths and make the overall dust temperatures lower. There is evidence for millimeter-sized dust grains in protostellar envelopes which would cause a shallower \(\beta\) of \(~1\) (Sadavoy et al. 2013; Kwon et al. 2009; Schnee et al. 2014) and we therefore feel justified in adopting a hybrid dust opacity model.

We generated 30,000 AU \(\times\) 30,000 AU model images with 15 AU resolution, corresponding to 2048 \(\times\) 2048 pixel images for emission between 2.8 mm and 3.0 mm. Such high resolution was necessary to ensure that we did not introduce false structure when Fourier transforming the images to compare with the observed visibility data. We used the MIRIAD task \textit{fft} to calculate the Fourier transform of each model image and we azimuthally averaged the Fourier transformed image to construct a one-dimensional visibility amplitude profile. To facilitate model comparisons, we normalized the flux densities of the envelopes and data in our comparisons at \(uv\) distances of 22.5 \(k\lambda\). This normalization is reasonable because most emission should be optically thin at \(\lambda \sim 3\) mm on spatial scales larger than our best resolution (\(~2000\) AU) and we are only interested in comparing the density profiles, not fitting model envelopes in detail.

We perform a simple model comparison for the sources with the best signal-to-noise ratios at the \(uv\) distances probed by our observations: 082005, 093005, 090003, 082012, and 097002, see Figures 5 and 6. This sample is representative of the range in visibility amplitudes profiles observed, e.g., from very flat (090003, 097002, and 093005), intermediate (082005), and rapidly declining (082012).

The flat visibility amplitude sources (090003, 097002, and 093005) are consistent with a power-law envelope with \(\rho \propto R^{-2.5}\) if there is no unresolved component. When a 0.01 \(M_\odot\) unresolved component is included in a power-law envelope, the flat visibility sources are still most consistent with a \(\rho \propto R^{-2.5}\) envelope. We note that the observed visibility amplitude profiles of 090003, 097002, and 093005 are systematically elevated with respect to the \(\rho \propto R^{-2.5}\) envelope both with and without the inclusion of a 0.01 \(M_\odot\) disk. When comparing the flat visibility sources to models with a 0.1 \(M_\odot\) unresolved component, all three power-law envelope density models provide a reasonable match to the visibility data since the compact source dominates the emission at \(uv\) distances \(>20\,k\lambda\), although the
curvature of visibility curve of the $R^{-1.5}$ model is opposite to that apparent in the data. In the case of the CMU envelope, the flat visibility amplitude sources are inconsistent with no disk component and are consistent with a 0.1 $M_\odot$ disk with $R_C \leq 100$ AU, but the curvature of the model visibility amplitude profile is in the opposite sense as the data.

The intermediate source between the extremes of flat visibilities and rapidly declining visibilities (082005) is consistent with a density profile between $\rho \propto R^{-2.0}$ and $\rho \propto R^{-1.5}$ if no unresolved component is included. It is consistent with $\rho \propto R^{-2.0}$ when a 0.01 $M_\odot$ unresolved component is included, but is marginally inconsistent with a power-law envelope and a 0.1 $M_\odot$ unresolved component. Assuming a CMU envelope, it is most consistent with $R_C = 300$ AU and a 0.1 $M_\odot$ disk.

The rapidly declining visibility amplitude source 082012 is well-matched by the power-law envelope models with no unresolved component and a density profile between $\rho \propto R^{-2.0}$ and $\rho \propto R^{-1.5}$. The data are also consistent with $\rho \propto R^{-1.5}$ when a 0.01 $M_\odot$ disk is included. If we assume a CMU envelope structure, then 082012 is also consistent with a CMU envelope with $R_C = 100$–300 AU, containing a 0.01 $M_\odot$ disk. Thus, sources with rapidly declining visibility amplitudes (082012 and others in the sample) are inconsistent with both density profiles steeper than $\rho \propto R^{-2}$ and disk components more massive than 0.01 $M_\odot$.

In addition to comparing the visibility amplitude profiles directly, we show the visibility amplitude ratios from the envelope-only models in Figure 4. These ratios can be thought of as limiting cases in the absence of a massive protostellar disk. The visibility amplitude ratio for the $\rho \propto R^{-2.5}$ envelope is the smallest, but it is still in excess of observed sources with the smallest ratios. Most observed sources have visibility amplitude ratios in between the values found for the $\rho \propto R^{-2.5}$ and $\rho \propto R^{-2.0}$ models. The addition of an unresolved component to any of the models in Figure 4 would decrease the ratios, making the models more consistent with the observations, but with a shallower density profile.

The qualitative model comparison shows that multiple physical structures can be invoked to explain both the flat and rapidly declining visibility amplitude sources. The flat visibility sources can be explained with having most flux in the unresolved component or a very steep ($\rho \propto R^{-2.5}$) density profile. Thus, a steep density profile is essentially indistinguishable from a compact source with our current data. The rapidly declining visibility sources, on the other hand, are inconsistent with a massive unresolved component (0.1 $M_\odot$) within a power-law or CMU envelope. There may, however, be some additional dependence on the disk density structure that we do not explore here. Higher resolution data will be necessary to break these degeneracies between power-law envelopes with no or an unresolved component and rotationally flattened envelopes with a large, massive disk component.

4. DISCUSSION

The PBRS represent an intriguing piece to the puzzle of low-mass star formation. Their 2.9 mm luminosities are quite large relative to their bolometric luminosities, and their millimeter luminosities are comparable to the brightest millimeter sources known in the nearby star forming regions. At the same time 4 out of the 14 observed PBRS have some of the flattest 2.9 mm visibility amplitudes observed toward any protostellar source; indeed, the most comparably flat source is NGC 1333 IRAS 4B (Looney et al. 2003). In the following subsections, we compare and contrast the PBRS to known protostars in nearby star forming regions and theoretical models to examine their significance in the star formation process as a whole.

4.1. Envelope Density Profiles

The favored interpretation of the very red 24–70$\mu$m colors exhibited by the PBRS is very high envelope densities (ST13). Most of the observed PBRS envelopes appear to be quite massive as measured from their 2.9 mm flux densities (Table 2). Furthermore, the low 5–30$\mu$m flux ratios indicate that there is
Figure 5. Plots of normalized visibility amplitudes vs. uv distance for models and a few representative sources overlaid. We plot radial density profiles ($\rho \propto R^{-1.5}$, $-2$, $-2.5$) and compact object (disk) masses of $0.0 \, M_\odot$, $0.01 \, M_\odot$, and $0.1 \, M_\odot$ in each plot. The flux densities are normalized at $22.5 \, k\lambda$ ($9.5$; $3850 \, AU$) to limit the contribution of flux external to the envelopes and only probe internal structure. With no compact structure (i.e., no disk), the sources with flat visibility amplitudes are consistent with either having a density profile as steep as $\rho \propto R^{-2.5}$ or being dominated by unresolved structure, while 082012 is inconsistent with being dominated by unresolved structure.

A singular isothermal sphere (SIS) also has a $\rho \propto R^{-2}$ density profile and this is the initial condition of the (Shu 1977) protostellar collapse model. The free-fall collapse of a SIS is inside-out, meaning there is an outwardly propagating rarefaction wave that bounds the infalling region of the envelope; the infalling region has a $\rho \propto R^{-1.5}$ density profile. This model was extended to include rotation by Terebey et al. (1984) and the region inside the centrifugally supported radius has a density profile of $\rho \propto R^{-0.5}$. Within the context of these models, the envelope emission of the youngest sources is expected to be dominated by the $\rho \propto R^{-2}$ region since both the infall and rotationally supported regions are small at early times. For an initial sound speed ($c_s$) of $0.2 \, km \, s^{-1}$ (assuming $T=10 \, K$), the infalling region would extend to a radius of $\sim 1050 \, AU$ ($2.5$) $25 \, kyr$ after collapse begins ($r = c_s \times t$). Thus, for extremely young sources, one would expect that data with a maximum resolution of $\sim 2''$ to be dominated by the $\rho \propto R^{-2}$ component of the envelope (within the context of this model). If some of the sources, however, are more evolved, with larger collapsing regions, the shallower density profiles and possibly

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the rotationally flattened portion of the density structure would be apparent in the visibility amplitudes (see Figure 6).

The envelope models that we compare to the observed visibility amplitudes in Figure 5 have density profiles that encompass those expected from the theories of protostellar collapse. Considering only the envelope contribution to the visibility amplitudes without additional unresolved source emission, however, the sources with flat visibility amplitudes (090003, 093005, 097002, 135003 and 091016) models are most consistent with the envelope density profiles that are as steep as $\rho \propto R^{-2.5}$.

The sources with more rapidly declining visibility amplitudes (082012 and 302002) are consistent with the $\rho \propto R^{-1.5}$ density profile. The sources in between (082005 and 091015) are consistent with $\rho \propto R^{-2}$ density profiles. Thus, if we consider only envelopes with smooth radial density profiles, the flat visibility amplitude sources appear to be inconsistent with the analytic protostellar collapse theories, while the sources with intermediate and rapidly declining visibility amplitudes fall within theoretical expectations.

Such steep density profiles have been obtained previously toward more nearby protostars (Looney et al. 2003; Chiang et al. 2008, 2012; Kwon et al. 2009). The sources with flat visibility amplitudes appear similar to NGC 1333 IRAS 4A or IRAS 2A. Our results further confirm that the dust continuum structure of some protostellar envelopes indicate radial density profiles steeper than expected from analytic models of collapse, consistent with previous findings by Looney et al. (2003) and Kwon et al. (2009).

The density profiles steeper than the analytic models could result from having asymmetric envelope structures on these scales. Tobin et al. (2010, 2012) showed that envelopes are often filamentary and asymmetric on $>1000$ AU scales and that infall might come through a filamentary envelope rather than a spherical envelope. The Spitzer IRAC and the APEX 350 $\mu$m and 870 $\mu$m images in ST13 often show filamentary structure on larger scales that may persist on smaller scales (e.g., Figures 8 and 13(b) of ST13). The 2.9 mm continuum images described here, however, do not have features suggestive of strong asymmetry in most cases, but Tobin et al. (2010) argued that asymmetry can be difficult to observe in the dust continuum due to the emission resulting from a combination of density and temperature.

If additional, unresolved components to the dust emission are considered (i.e., added to the envelope density profile), the flat visibility amplitude sources may be consistent with shallower density profiles. With an unresolved component of $0.1 M_\odot$, visibility amplitudes out to $\sim 150$ kA, at the distance to Orion. The other sources with more rapidly declining visibility amplitudes appear more similar to NGC 1333 IRAS 4A or IRAS 2A. Our results further confirm that the dust continuum structure of some protostellar envelopes indicate radial density profiles steeper than expected from analytic models of collapse, consistent with previous findings by Looney et al. (2003) and Kwon et al. (2009).
then the $\rho \propto R^{-2.0}$ and $R^{-1.5}$ density profiles (in addition to $\rho \propto R^{-2.5}$) are able to reproduce the flat visibility amplitudes observed for some sources (Figure 5). Even with the unresolved component, however, the curvature in the visibility amplitudes of the $\rho \propto R^{-1.5}$ density profile is inconsistent with the data. The steeper density profiles also do not fully capture the curvature observed in the data, but the effect is less dramatic than for $\rho \propto R^{-1.5}$.

A circumstellar disk is a naturel compact structure that is expected to develop during the star formation process (e.g., Williams & Cieza 2011, and references therein) and are likely to have a variety of sizes (Maury et al. 2010; Tobin et al. 2012). Therefore, a comparison to the CMU models with and without disk components is also of interest, see Figure 6). The rapidly declining and intermediate sources could be consistent with having a large $R_C$ region (and a large protostellar disk), depending on the disk mass. The sources with flat visibility amplitudes can only be consistent with small $R_C$ and a massive disk. Again, however, the curvature of the model visibility amplitude curves with small ($R = 50 \, \text{AU}$, 100 AU), $0.1 \, M_\odot$ disks is dissimilar to that of the data. Thus, the models show that disk components are possible in the intermediate and rapidly declining cases, but the exact parameters are not well-constrained. The flat visibility amplitude sources on the other hand are grossly consistent with a disk component within the context of the CMU model, but, as stated previously, the curvature of the visibility amplitude profiles of the models relative to the observations are different.

Rather than a massive disk component, it is also possible that the envelope density profile itself is not a smooth power law. Simulations of protostellar collapse including magnetic fields have been shown to create density enhancements of infalling material during collapse that depart from a smooth power-law density profile (Tassis & Mouschovias 2005). Such structures can cause flattening of the visibility amplitudes due to the mass build-up at small spatial scales. The robustness of this model, however, is unclear.

Regardless of the exact nature of the structure, we can conclude that the envelopes with flat visibility amplitudes are inconsistent with the often assumed $\rho \propto R^{-1.5}$ density profile for protostellar envelopes. The flat visibility amplitude profiles are most consistent with either a steep density profile ($\rho \propto R^{-2.5}$) or the visibility amplitudes are dominated by dense, unresolved structure. The unresolved structure could be a disk, or it could be departures from a power-law density profile. Higher resolution data that resolve down to the expected scales of a disk are necessary to distinguish between these two (radically different) scenarios.

4.2. Nature of the PBRS

The defining characteristic of the PBRS from the Herschel study by ST13 is the very red color of the PBRS sample as a whole, having $[24 \, \text{µm} ]$–$[70 \, \text{µm} ]$ colors (in $\log (\lambda F_\lambda)$ space) redder than 1.65. Furthermore, the $T_{\text{bol}}$ measurements of the PBRS are in a narrow range of $20$–$45 \, \text{K}$, consistent with little observed emission shortward of $24 \, \text{µm}$ for most PBRS. Thus, while the coldness (and redness) of the SEDs of the PBRS is consistent throughout the sample, they are very heterogeneous in terms of their ratio of $L_{\text{sub-mm}}$ to $L_{\text{bol}}$ (0.6%–6.1%), $L_{\text{bol}}$ itself, and 2.9 mm luminosity. Nevertheless, the PBRS are characterized by higher $L_{2.9 \, \text{mm}}$ to $L_{\text{bol}}$ ratios than previously identified in protostellar samples (see Figure 2).

The expected evolutionary trend for protostars is that they become more luminous as they accrete mass due to increased photospheric luminosity and greater accretion luminosity (Young & Evans 2005; Dunham et al. 2010). $T_{\text{bol}}$ is also expected to also increase with decreasing envelope density and hence optical depth. At the same time, clearing of the envelope is likely driven by the influence of the protostellar outflow (Arce & Sargent 2006; Offner & Arce 2014). Thus, it is expected that very young Class 0 protostars will have larger millimeter flux densities (or larger envelope mass), with rather low luminosities; in other words they will have large fractions of millimeter luminosity relative to $L_{\text{bol}}$ (e.g., Andre et al. 1993). These changes, however, may not be due solely to evolution: both initial conditions and evolution play roles in the observed $T_{\text{bol}}$ and $L_{\text{sub-mm}}/L_{\text{bol}}$ ratios observed toward particular protostars (e.g., Young & Evans 2005). Further complicating matters, $T_{\text{bol}}$ can be strongly influenced by the orientation of a given source in the plane of the sky, such that a more evolved source viewed edge-on can appear younger (ST13; Jørgensen et al. 2009; Launhardt et al. 2013; Dunham et al. 2014). Within the sample of PBRS, $T_{\text{bol}}$ is confined to a narrow range, but there are a few sources that have low luminosities and low 2.9 mm flux densities (061012 and 119019). Meanwhile, other PBRS have both relatively high luminosities and high millimeter flux densities (082012 and HOPS 373). Such variations in the observed properties of the sample indicate that, despite the stringent color selection (ST13), the PBRS as a whole may not be characterized by a single evolutionary state. Furthermore, the fact that the PBRS present such a narrow range in $T_{\text{bol}}$ but exhibit a broad range in other properties poses possible problems for a $T_{\text{bol}}$-based classification of protostars. If even at the lowest values of $T_{\text{bol}}$ and for a uniformly selected sample we see clear variations, then even larger variations may be seen across the $T_{\text{bol}}$ range encompassing the Class 0 and Class I protostellar phases. Thus, not all protostars of equivalent $T_{\text{bol}}$ are equal.

With the interferometry data, we are able to determine the spatial scales from which we are detecting emission due to the high-resolution and analysis of visibility amplitude profiles. The visibility amplitude profiles of the PBRS have many features, but they can be broadly described as flat or rapidly declining. The visibility amplitude ratios from 5 k$\lambda$ to 30 k$\lambda$ (flux at $\sim$17,000–3000 AU scales) plotted versus 30 k$\lambda$ flux density and $L_{\text{bol}}$ (see Figure 4), further enable the sample to be examined as a whole. The expected evolutionary trends for visibility amplitude ratios and profiles are uncertain due to the unknown contribution of the disk at a given time. Nevertheless, we can use the analytic models for protostellar envelopes as limiting cases.

If we consider the collapse of an SIS with a density profile $\rho \propto R^{-2}$, the visibility amplitude ratios would be smaller initially and then increase as material falls in from larger radii. The density profile of the infalling region will be proportional to $R^{-1.5}$ and this region grows with time. So, in the absence of a massive disk, the visibility amplitude ratio for such a model will be between the $R^{-2}$ and $R^{-1.5}$ values (see ratios taken from the models in Figure 4). Then as a disk grows and the envelope dissipates, the 5–30 k$\lambda$ visibility amplitude ratios will decrease (trending toward 1 on the scales examined) as the disk begins to dominate the flux density of the system at millimeter wavelengths. Thus, to summarize the visibility amplitude ratios will start small, then increase, and then decrease. The exact evolution will depend on how quickly a large disk forms and how massive such a disk is. We note that if an SIS formed from a Bonnor–Ebert sphere, a Bonnor–Ebert sphere would initially
have a large visibility amplitude ratio, depending on the size of the flat density region (e.g., Schnee et al. 2012). The visibility amplitude ratios would then decrease as the Bonnor–Ebert sphere evolved toward an SIS.

Two trends are evident in the visibility amplitude ratio plots shown in Figure 4. Most of the sources with the highest flux densities at 30 kλ also have the lowest 5–30 kλ ratios (i.e., spatially compact flux density). Then looking at the 5–30 kλ ratio versus $L_{\text{bol}}$ we see that 6 out of 14 of the lower luminosity PBRS sources ($L_{\text{bol}} < 3 L_\odot$) have ratios between 1 and 1.5. The non-PBRS and higher luminosity PBRS tend to have higher ratios, except for the sources that are the most evolved (i.e., HOPS 223 and HOPS 59). We therefore suggest that the visibility amplitude ratios enable us to divide the PBRS sample into two groups. The PBRS with the smallest visibility amplitude ratios (flattest profiles) are the youngest of the PBRS and the sources with larger ratios (rapidly declining amplitudes) are likely more evolved, though still Class 0 protostars.

The youngest sources would then have the most compact, dense envelopes initially that may then be accreted rapidly due to the short free-fall time. The flat visibility amplitudes indicate that there is likely large amounts of mass within only a few thousand AU, implying high average densities. A source with $2 M_\odot$ of envelope material a radius of 1500 AU has an average density of $3 \times 10^7$ cm$^{-3}$, corresponding to a local free-fall time of $\sim$10 kyr. Thus, a substantial amount of the final protostellar mass could be accumulated in this short period of time, much less than the expected lifetime of the Class 0 phase ($\sim$150 kyr) (Dunham et al. 2014). Therefore, the free-fall times suggest that the Class 0 phase may begin with a short period of rapid infall that may only last $\sim$10% of the Class 0 phase. Based on the number of detected sources, ST13 suggested that if the PBRS represented a phase of protostellar evolution distinct from the Class 0 phase, it may only last $\sim$25 kyr. Rather than necessarily being a distinct phase, we believe that the PBRS with flat visibility amplitudes (093005, 090003, 091015, 091016, 082005, and 097002) are among the youngest Class 0 protostars, and the large amount of mass on small spatial scales could indicate that they are in a brief period of high-infall/accretion.

The PBRS that do not have flat visibility amplitudes are still young, but may be more comparable to typical Class 0 sources. We suggest that the sources with bright 2.9 mm flux densities, but rapidly declining visibility amplitudes (302002, 082012, HOPS 373) are slightly more evolved that the PBRS with flat visibility amplitudes. At least a fraction of their inner envelopes is likely to have been accreted onto the disk and/or protostar. The remaining sources with declining visibility amplitudes and low flux densities (119012, 061012, HOPS 372, 135003, 019003) are still consistent with being young Class 0 sources. Their cold $T_{\text{bol}}$ values and extremely red 24–70 μm, however, colors could result from high density envelopes, but with less overall mass. Alternatively, they could be edge-on sources; we will further explore the properties of these sources in relation to their outflows in an upcoming paper (J. J. Tobin et al., in preparation).

If the proposed scenario is true, then the Class 0 phase might be a two-phase process, with a short, rapid accretion phase (like a Bonnor–Ebert collapse; see Foster & Chevalier 1993)), lasting $\sim$10–25 kyr. This phase is then followed by a period of slower mass assembly, for the remainder of the Class 0 phase ($\sim$100–150 kyr), assuming a Class 0 lifetime of $\sim$160 kyr (Dunham et al. 2014). This idea is consistent with the models of Offner & Arce (2014) that show most protostellar mass being accreted during the Class 0 phase, before the outflow destroys the envelope.

### 4.3. Comparison to V eLLOs and Candidate FSHCs

The infrared and millimeter properties of the PBRS distinguish them from typical Class 0 protostars and indicate that at least some of the PBRS may be very young Class 0 objects in a period of high infall. Two other sub-classes of protostars identified by Spitzer and submm/millimeter observations are the V eLLOs and candidate FSHCs and it is important to distinguish the PBRS from these sources based on their millimeter properties. First, many of the V eLLOs and candidate FSHCs are very faint at 2.9 mm. For instance, the brightest V eLLO/candidate FSHC at 2.9 mm is Per-Bolo 58 with a flux density of 13 mJy at $d \sim$ 230 pc (Schnee et al. 2010; Enoch et al. 2010). If this source was at the distance to Orion, it would have a flux density of only $\sim$3.9 mJy and only appear as a $\sim$4σ detection in our data. This flux density is less than half that of the faintest source in the PBRS sample (119019). None of the other V eLLOs or FSHC candidates would be detectable at the distance to Orion with the sensitivity of our CARMA observations. The PBRS 061012 may be the most similar to a V eLLO, having the lowest luminosity, but it has a higher 2.9 mm flux density than other V eLLOs.

A comparison to the visibility amplitudes of the V eLLOs is less straightforward. Most V eLLOs/candidate FSHCs are not bright enough to enable analysis of their visibility amplitude profiles. Per-Bolo-58 is found to have rapidly declining visibility amplitudes and The candidate FSHC L1451-MMS (Pineda et al. 2011), however, has flat visibility amplitudes at 1.3 mm, similar to some PBRS sources. The visibility amplitudes are 30 mJy out to $\sim$200 kλ or 230 AU scales. At 2.9 mm, the visibility amplitudes of this source would be 2.5 mJy, assuming $\beta = 1$. At the distance to Orion, however, the visibility amplitudes would be below our detection limits at 0.8 mJy. While the overall emitting mass of this source is much lower than the PBRS, it does have a similar 5–30 kλ flux ratio, meaning that the envelope density profile might be similar to the most concentrated PBRS. However, L1451-MMS is undetected at 70 μm and 100 μm, unlike the PBRS. In summary, the millimeter properties of the PBRS, combined with the far-infrared constraints from ST13, distinguish the PBRS from the V eLLOs and candidate FSHCs.

### 5. SUMMARY AND CONCLUSIONS

We have presented CARMA 2.9 mm dust continuum observations toward 14 PBRS (Stutz et al. 2013) in the Orion A and B star forming regions, 12 of these protostars were first identified by Herschel observations. This sample of 14 PBRS also includes 135003, a new PBRS that discovered by Herschel and was not included in the Stutz et al. (2013) sample due to their stringent FWHM cut-off. The inclusion of this source increases the total number of PBRS in Orion to 19. The PBRS classification in Stutz et al. (2013) required $[24\mu m]–[70\mu m] > 1.65$. In addition, we also report the continuum properties of four protostars and one apparent starless/pre-stellar core within the fields observed toward the PBRS.

All 14 PBRS are detected in dust continuum emission. Twelve out of 14 have flux densities $>30$ mJy and three have flux densities $>90$ mJy. The eight PBRS with $L_{\text{bol}} \sim 1 L_\odot$ exhibit higher 2.9 mm luminosities than other known
protostars with similar \( L_{\text{bol}} \) values, and therefore have characteristics not previously identified. The PBRS with \( L_{\text{bol}} > 2.7 L_\odot \) have comparable 2.9 mm luminosities yet lower \( L_{\text{bol}} \) values than the brightest sources in the more nearby regions of Perseus and Ophiuchus. Furthermore, the 2.9 mm luminosity does not strongly correlate with \( L_{\text{bol}} \). This lack of correlation indicates that the 2.9 mm luminosity is not strongly dependent on either the central protostellar luminosity or the accretion luminosity.

Six PBRS sources (097002, 090003, 093005, 082005, 091015, and 091016) have flat 2.9 mm visibility amplitudes (and \( 5-30\,\text{k}\)) visibility amplitude ratios of less than 2). As a consequence, more than \(~50\%\) of the total flux density arises from scales that are smaller than \( 7'' (\sim 3000\,\text{AU}) \) in diameter. This behavior indicates either steep envelope density profiles or the presence of significant mass contained within a compact, unresolved structure. We suggest that these particular PBRS are the youngest of the sample and may be in a brief period of high infall rate. Indeed, the average density on scales \(<3000\,\text{AU}\) implies local free-fall times of \( \sim 10\,\text{kyr} \), in agreement with independent life-time estimates based on the ratio of PBRS to protostars (Stutz et al. 2013). The PBRS with large 2.9 mm flux densities but rapidly declining visibility amplitudes (302002, 082012, HOPS 372) are still considered to be young Class 0 protostars, but may be more evolved than the PBRS with flat visibility amplitudes. The sources with lower 2.9 mm flux densities and declining visibility amplitudes (119012, 061012, HOPS 372, 135003, 019003) are also still consistent with being Class 0 sources, but may have edge-on orientations and/or lower envelope masses.

To better characterize the density profiles of the sample, we compare the observed visibility amplitudes of the sources to Hyperion radiative transfer models of axisymmetric envelopes with varying radial density profiles and unresolved components (represented by a disk component). We also compare with rotating collapse models with various centrifugal radii and disk masses. We find that without an unresolved component to the emission, the flat visibility amplitude sources are most consistent with a \( \rho \propto R^{-2.5} \) radial density profile. If a compact structure is massive enough, however, then all three envelope density profiles tested here (\( \rho \propto R^{-1.5, -2.0, -2.5} \)) are able to provide a reasonable match to the data. Thus, with the current data we cannot distinguish between these two scenarios and higher resolution data are required to understand the nature of these sources. Furthermore, sources with more rapidly decreasing visibility amplitudes may be consistent with shallower density profiles and are inconsistent with having a massive unresolved component.

While the PBRS occupy a narrow range to \( T_{\text{bol}} \), their 2.9 mm flux densities and visibility amplitude profiles show a large amount of heterogeneity suggesting that they are not all in exactly the same evolutionary stage. We suggest an evolutionary trend in which the sources with flat visibility amplitude profiles are the youngest and perhaps have a dense inner envelope that may be rapidly accreted. The sources large 2.9 mm flux densities and rapidly declining visibility amplitudes may be slightly more evolved than those with flat visibility amplitudes. Moreover, the sources with flat visibility amplitude profiles also tend to have lower \( L_{\text{bol}} \) values than those with rapidly declining visibility amplitudes, consistent with the expected evolutionary trend of increasing \( L_{\text{bol}} \). The PBRS also draw a sharp contrast with candidate FHSCs and VeLLOs. The PBRS have higher \( L_{\text{bol}} \) and larger 2.9 mm luminosities than all of the candidate FHSCs and VeLLOs; at the distance to Orion, none of the known candidate FHSCs or VeLLOs would have been confidently detected.

In summary, the PBRS have properties that are consistent with placing them among the youngest Class 0 protostars. Their millimeter and infrared properties distinguish them from typical Class 0 protostars, as well as from candidate FHSCs and VeLLOs. While the data presented here have enabled us to postulate a tentative evolutionary scenario, further characterization at higher resolutions and in molecular lines will be necessary to more firmly establish their place in the context of the star formation process.

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**APPENDIX**

**NOTES ON INDIVIDUAL SOURCES**

*A.1. 135003*

The PBRS source 135003 is located in the OMC2/3 region of the Orion A cloud. This source was not originally included in the PBRS sample of ST13, due to it having a 70 \( \mu \)m FWHM size slightly greater than 7.8. A multi-wavelength plot of this source and its SED are given Figures A1 and A2. The 24–70 \( \mu \)m color (\( \log(\lambda F_\lambda(70 \mu m)/\lambda F_\lambda(24 \mu m)) \)) for 135003 is 2.44 and it has \( L_{\text{bol}} = 11.2 L_\odot \), \( T_{\text{bol}} = 34.6 \, \text{K} \), \( L_{\text{sub-mm}}/L_{\text{bol}} = 0.175 \), and an SED peak at \( \sim 121 \mu \)m. These values are derived from modified blackbody fitting, the same method used in ST13. The photometry for 135003 is given in Table A1. The brighter, neighboring source is HOPS 59, a Class I protostar and separated by 18′′ (\( \sim 7600 \, \text{AU} \)); HOPS 59 is bright both at 2.9 mm and in the infrared. There is another 3σ source located between 135003 and HOPS 59, which may be a real source given that there is an IRAC point source detected at this position at 3.6 \( \mu \)m and 4.5 \( \mu \)m. The northeastern extension of emission from 135003 is real and has been observed in 3 mm single-dish maps of the region (Schnee et al. 2014), in addition to the 350 \( \mu \)m and 870 \( \mu \)m maps showing similarly extended structure. The visibility amplitudes of this source appear to flatten toward the longer baselines, while there is some up and down variation likely due to the extended emission and surrounding sources.
### Table A1
135003 Photometry

| Wavelength (μm) | Flux Density (Jy) |
|----------------|------------------|
| 24.0           | <0.016           |
| 70.0           | 13.1 ± 0.7       |
| 100.0          | 32.9 ± 3.3       |
| 160.0          | 88.9 ± 6.2       |
| 350.0          | 11.4 ± 2.2       |
| 870.0          | <2.8             |

Nevertheless, the overall flux density at large uv distances is less than other PBRS.

#### A.2. 093005

The reddest PBRS in 24–70 μm color, 093005, is located at the intersection of three filaments and ∼110′′ north of its nearby neighbor of HOPS 373 (ST13). At wavelengths shorter than 70 μm, 093005 was only detected in Spitzer 4.5 μm imaging (ST13). The dust continuum emission from this source is quite compact, only ∼1.2× fainter at 30 kλ versus 5 kλ and the visibility amplitudes are very flat out to ∼80 kλ.

#### A.3. HOPS 373

HOPS 373 is the close neighbor of 093005, located 110′′ to the south. The dust continuum emission observed in D-configuration only showed some asymmetry and the combined D and C configuration data resolve a second component, separated by 4′′. The continuum emission from HOPS 373 is fainter than 093005 and the visibility amplitudes fall quickly with increasing uv distance, reaching zero at 50 kλ and then rise again, showing the expected pattern for two sources separated by 4′′.

#### A.4. 090003

The PBRS 090003 (also called Orion B9 SMM 3; Miettinen et al. 2012) is located in a loose filamentary complex with several protostars and starless cores over a 0.5 pc region (Miettinen et al. 2012). Much like 093005, its only detection shortward of 24 μm is at 4.5 μm, where there is a small feature offset from the location of the protostar, indicative of shocked H₂ emission (Miettinen et al. 2012; Stutz et al. 2013).

The envelope as observed at 870 μm is extended to the east and at 350 μm there are a few sub-peaks associated with the extended 870 μm emission (Miettinen et al. 2012). We do not, however, detect 2.9 mm continuum emission from these.
sub-peaks, indicating that they do not harbor compact protostellar sources. The continuum emission from 090003 has a very similar visibility amplitude profile to 093005, indicating that this source is also dominated by small-scale emission. The source 090003 is the second brightest in the sample at 2.9 mm.

A.5. 082012 and HOPS 372

The brightest continuum source in the sample is 082012; the 2.9 mm continuum luminosity of 082012 is about the same as that of the well-studied multiple protostellar system IRAS 16293-2422. Similarly, 082012 has visibility amplitudes that fall rapidly as a function of $\lambda$, indicating that the source is not dominated by compact emission. Moreover, 082012 ($L_{bol} = 6.3 L_\odot$) has a lower luminosity than other protostars with similar envelope characteristics. Only $\sim 20''$ away from 082012 is HOPS 372, another PBRS source, but with much lower overall continuum flux.

A.6. 019003

The PBRS source 019003 is also located in the OMC 2/3 region, northward of 135003. We detect four sources at 2.9 mm surrounding 019003 (Figure 1); HOPS 71 farthest to the east, HOPS 68 to the south, and 019003 itself, which resolves into two sources denoted 019003 A and 019003 B. The source 019003 A is most closely associated with the 70 $\mu$m source detected by ST13. The 70 $\mu$m source, however, is offset by $\sim 3''$ to the west, as are the 4.5 $\mu$m and 24 $\mu$m sources. On the other hand, 019003 B appears to be starless, with no detections at any wavelength shorter than 160 $\mu$m.

A.7. HOPS 68

The protostar HOPS 68 is detected at the edge of the 019003 field. This source was previously recognized for the appearance of crystalline silicates in its Spitzer IRS spectrum (Poteet et al. 2011).

A.8. 302002

The source 302002 is located at the end of an isolated filamentary structure in NGC 2068, $\sim 20''$ to the south of a Class I source (HOPS 331). The 2.9 mm continuum is well-detected toward 302002, whereas HOPS 331 is undetected. The visibility amplitudes toward this source are falling with increasing $\nu$ distance, similar to those of 082012.

A.9. 061012 and HOPS 223

The source 061012 is located very near the outbursting protostar V2775 Ori (HOPS 223; also detected at 2.9 mm) in the L1641 region (Fischer et al. 2012). This source is one of the faintest PBRS in the sample at 17 mJy; indeed, only 119019 is slightly fainter. The visibility amplitudes appear rather flat, but become dominated by noise at $\nu$ distances larger than 20 $\kappa\lambda$. Hence, its nature is uncertain.

A.10. 091015 and 091016

The two sources 091015 and 091016 are close neighbors in NGC 2068, the former being $\sim 40''$ west of the latter. These sources are completely undetected at wavelengths shortward of 70 $\mu$m. The visibility amplitudes are flat across all observed $\nu$ distances toward both sources in Figure 3, with 091015 declining a bit more rapidly than 091016.

A.11. 082005

The source 082005 is located about 4' south of 082012, and these sources are connected by an apparent filamentary structure seen at 160 $\mu$m and 870 $\mu$m. This source is also undetected at wavelengths shorter than 70 $\mu$m, like 091015 and 091016. The visibility amplitude data for the dust continuum are rather flat like some of the other deeply embedded sources, but still decline more rapidly than the flattest sources.

A.12. 097002

The source 097002 is found near a bright 4.5 $\mu$m and 24 $\mu$m source seen in Spitzer data, as shown by ST13. At 70 $\mu$m, the neighboring 24 $\mu$m source is undetected, but 097002 is quite apparent. We detect 097002 as a bright 2.9 mm continuum source with very flat visibility amplitudes, but we do not detect the neighboring source. Much like 090003 and 093005, the visibility amplitudes toward this source are very flat.

A.13. 119019

The source 119019 is the faintest PBRS source in terms of 2.9 mm continuum emission and appears more similar to the protostars observed in the Tobin et al. (2011) sample rather than the PBRS sources.

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