The globular cluster system of the young elliptical NGC 6702

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\textbf{ABSTRACT}

We study the globular cluster (GC) system of the dust-lane elliptical galaxy NGC 6702, using $B$, $V$ and $I$-band photometric observations carried out at the Keck telescope. This galaxy has a spectroscopic age of $\approx 2$ Gyrs suggesting recent star-formation. We find strong evidence for a bimodal GC colour distribution, with the blue peak having a colour similar to that of the Galactic halo GCs. Assuming that the blue GCs are indeed old and metal–poor, we estimate an age of 2–5 Gyrs and supersolar metallicity for the red GC subpopulation. Despite the large uncertainties, this is in reasonable agreement with the spectroscopic galaxy age. Additionally, we estimate a specific frequency of $S_N = 2.3 \pm 1.1$ for NGC 6702. We predict that passive evolution of NGC 6702 will further increase its specific frequency to $S_N \approx 2.7$ within 10 Gyrs, in closer agreement to that of typical present-day ellipticals. We also discuss evidence that the merger/accretion event that took place a few Gyrs ago involved a high gas fraction.

\textbf{Key words:} globular clusters: general – galaxies: individual: NGC 6702 – galaxies: star clusters.

\section{1 INTRODUCTION}

The ‘merger’ hypothesis, postulating that elliptical galaxies form by spiral galaxy mergers, was first proposed by Toomre & Toomre (1972). This scenario has gained new impetus over the last few years by both numerical simulations (e.g. Barnes 1992; Hernquist 1992; Mihos & Hernquist 1996) and observational studies, suggesting that the properties of merger remnants are consistent with those of ellipticals (e.g. Toomre 1977; Schweizer et al. 1990; Casoli et al. 1991; Schweizer & Seitzer 1992; Hibbard & van Gorkom 1996; Georgakakis, Forbes & Norris 2000).

However, one of the long standing drawbacks of the ‘merger’ hypothesis is the excess globular clusters (GCs) per unit starlight around ellipticals compared to spirals (van den Bergh 1984). To overcome this problem Schweizer (1987) and Ashman & Zepf (1992) suggested that the starburst activity usually experienced during gas-rich mergers will form new GCs and thus increase their number compared to their progenitors. Indeed, there is accumulating observational evidence supporting this scenario in on-going mergers (e.g. Lutz 1991; Whitmore et al. 1993; Whitmore & Schweizer 1995; Schweizer 1996; Miller et al. 1997; Zepf et al. 1999; Forbes & Hau 2000). However, the situation in elliptical galaxies (the eventual product of mergers) is less clear. Although many ellipticals have been observed to have bimodal colour distributions, when examined in detail their GC properties do not match those from the merger picture (Forbes et al. 1997). Outstanding questions include whether the newly formed GCs survive and whether the final GC system resembles that of typical (old) ellipticals. Probing the GC properties in intermediate age merger-remnants (1–5 Gyrs) may be of key importance for addressing these issues. Indeed, such young ellipticals may provide the link between on-going mergers and relaxed ellipticals.

For example, a sample of merger–remnants spanning a range of ages will allow study of the chronological development of the system’s specific frequency (measuring the number of globular clusters per unit starlight) in comparison with that of evolved ellipticals. Young ellipticals are also expected to have two distinct GC populations with different colour and magnitude distributions. By exploiting these differences one can attempt to age-date the newly formed
GCs, which in turn can be compared to the time since the gaseous merger. Indeed, studies of the dynamically young proto-ellipticals NGC1700 and NGC3610 have shown the presence of bimodality in their colour distributions (White et al. 1997; Brown et al. 2000). Also the age of the new GCs, although model dependent, is found to be similar to estimates for the galaxy age based on spectroscopic or dynamical arguments.

In this paper we study the globular cluster system of the elliptical galaxy NGC6702. Lauer (1985) found a dust lane in the centre of NGC6702 indicating a recent gaseous merger event. This galaxy is also relatively luminous at far-infrared wavelengths compared to other elliptical galaxies, with $L_{FIR} = 1.3 \times 10^{10} L_\odot$ (Wiklind et al. 1995). Terlevich & Forbes (2000) have estimated a spectroscopic age and metallicity of about 2 Gyrs and $[Fe/H] = +0.5$ respectively for NGC6702 suggesting that it is a young elliptical. The main observational properties of NGC6702 are given in Table I.

Additionally, according to the NED database the nearest galaxy projected on the sky that is either brighter or no more than 2 magnitudes fainter, than NGC6702 is NGC6703 at a projected distance of 10 arcmin. However, as NGC6702 has a recession velocity of 4727 km s$^{-1}$ and NGC6703 has $V = 2365$ km s$^{-1}$ they are actually separated by some 30 Mpc assuming no peculiar motions. The nearest galaxy with a similar recession velocity is NGC6711 ($V = 4671$ km s$^{-1}$) with a separation of $\sim 120$ arcmin or over 2 Mpc. Other galaxies listed by NED are more than 2 magnitudes fainter than NGC6702. Thus NGC6702 is an example of an isolated elliptical galaxy with no comparable galaxies within 2 Mpc. Such galaxies are fairly rare and hence their evolutionary history is of some interest. In particular, ellipticals in low density environments are interesting within the hierarchical merger framework as they are expected to have assembled much more recently than their cluster counterparts (Kauffmann & Charlot 1998). Any merger would be expected to have occurred in the more recent past and involved progenitors that have largely processed gas into stars.

In section 2 the observations, data reduction and the isophotal properties of NGC6702 are presented. Section 3 describes the candidate GC selection, while section 4 presents our results which are discussed in section 5. Finally, section 6 summarises our conclusions. Throughout this paper we adopt $H_0 = 75$ km s$^{-1}$ Mpc$^{-1}$.

2 OBSERVATIONS AND DATA REDUCTION

2.1 Description of the data

Broad-band imaging of NGC6702 in the $B$, $V$ and $I$ filters was carried out at the Keck-II telescope on 1999 August 17, using the the Low Resolution Imaging Spectrometer (LRIS; Oke et al. 1995). The LRIS instrument, equipped with a TEK 2048 x 2048 CCD, is mounted on the Cassegrain focus providing a 0.215 arcsec pixel$^{-1}$ imaging scale and a 6$^\prime$ x 8$^\prime$ field–of–view. The total exposure time in the $B$, $V$ and $I$ filters was 3600, 1800, and 1200 sec respectively, segregated into three separate integrations of 1200, 600 and 400 sec respectively. The seeing at the time of the observations varied between 0.8 to 1 arcsec.

The data were reduced following standard procedures, using IRAF tasks. The reduced images are found to be flat to better than $\sim 2\%$. The final $V$-band image of NGC6702 is shown in Figure 1. Photometric calibration was performed using standard stars from Landolt (1992). The photometric accuracy, estimated using these standard stars, is $\pm 0.02$ mag in all three bands. We adopt a Galactic extinction in the NGC6702 direction of $E(B-V) = 0.06$, taken from Burstein & Heiles (1984). Assuming a standard reddening curve (Savage & Mathis 1979), we correct each magnitude for Galactic extinction using the values $A_B = 0.25$, $A_V = 0.19$ and $A_I = 0.09$ mag.

2.2 Galaxy surface brightness profile

The $B$-band surface brightness profile, $\mu_B$, of NGC6702 is plotted in Figure 2. This has been determined by fitting ellipses to the galaxy intensity profile using the ISOPHOTE task within STSDAS. During the ellipse fitting process the centre of the galaxy was kept fixed and a 3$\sigma$ clipping algorithm was employed. Also, bright objects as well as the central saturated parts of the galaxy, were masked out. The position angle (PA) and ellipticity ($\epsilon$) were fitted at each radius, until S/N constraints terminated the fitting process at low surface brightness. For NGC6702 this occurs at an equivalent radius $r_{eq} \approx 50^\prime$ ($r_{eq} = a \sqrt{1-\epsilon}$, where $a$ is the semi-major axis). Beyond that radius the position angle and ellipticity were kept constant. The sky level was determined using the method described by Goudfrooij et al. (1994).

Although the isophotes of elliptical galaxies are well approximated by ellipses, deviations at the few per cent level are common. Such deviations are often quantified by the 4th sine and cosine amplitudes ($S_4$ and $C_4$ respectively) of the Fourier transform of the galaxy intensity at a given radius. The standard interpretation of the residual $C_4$ terms is that
they indicate either the presence of a weak underlying disk if positive or boxiness if negative. However, it has been emphasised that both C4 and S4 terms should be considered when estimating the isophotal shape of galaxies (Franx et al. 1989; Peletier et al. 1990; Goodfrooij et al. 1994). For example, the presence of a disk, bar or ring with a significant position angle twist relative to the elliptical galaxy semi-major axis is expected to give both positive and negative C4 and/or S4 terms, depending on its orientation. Consequently, in the case of such projection effects we expect S4 terms to be significant.

Figure 3 plots the ellipticity, position angle as well as C4 and S4 terms as a function of radius for NGC 6702 in the B-band. Similar trends are apparent for the V and I-bands. The C4 term is negative at small galactocentric distances, indicating boxiness, i.e. excess light at 45° relative to the semi-major axis. This is partly due to the presence of the dust lane in the core of NGC 6702 (Lauer 1985) lying perpendicular to the semi-major axis and extending out to ≈ 20″. The S4 term is also non-zero within the same range of r_eq, albeit with a smaller amplitude. Masking out the dust lane reduces the amplitude of both the C4 and S4 terms. These terms are also non-zero (S4 positive and C4 negative) at r_eq ≈ 40″ indicating boxiness. At larger distances (r_eq > 60″) both the C4 and S4 deviate significantly from the elliptical symmetry, but this is likely to be an artifact arising from the low S/N at these radii.

2.3 Source extraction and completeness limits

To detect sources superimposed on NGC 6702 we subtract the galaxy light profile using a median filter technique. Firstly, the reduced image for each filter is smoothed using a median rectangular filter with a size of 40 pixels. The smoothed image is then subtracted from the original frame and the resulting image is subsequently used for source extraction and photometry.

The sources in the reduced frames were extracted using the SExtractor package (version 2.1.0; Bertin & Arnouts 1996). The main input parameters are the detection threshold, given as a multiple of the sky variance (σ_Sky) and the minimum area in pixels for an object to be extracted. After considerable experimentation, we adopted a threshold of 2.5 × σ_Sky and a minimum area of 10–15 pixels (≈ 0.5–0.7 arcsec²). This choice of values minimises the number of spurious detections, while ensuring that faint objects are successfully extracted. A total of 1455, 1456 and 1494 sources were extracted form the B, V and I frames respectively. The photometric catalogues generated for each waveband were then matched, resulting in a total of 939 common detections in all three filters, B, V and I. This source list is referred to as the matched catalogue. Using this catalogue we performed photometry using the PHOT task. After a curve–of–growth analysis on several sources an aperture of 8 pixels radius was employed with the background estimated from an annulus of 15 to 20 pixels around each source. For unresolved sources the 8 pixels radius aperture magnitude is close to total.

The completeness limit in a given filter is estimated by adding artificial point sources of different magnitudes to the original image, using MKOBJECTS task. Then we attempt to recover the artificial objects using SExtractor package, with the same parameter settings as with those adopted for the real data. A total of 1500 point sources were added to a given filter frame. To avoid overcrowding the frame, artificial sources are added in groups of 150 and the process is repeated 10 times. The results are presented in Figure 3, plotting the fraction of recovered objects as a function of magnitude in the B, V and I filters respectively. The 80% completeness limits in the B, V and I filters are found to be ≈ 26.1, ≈ 25.3 and ≈ 23.9 mag respectively. The mean photometric error at these magnitude limits is ≈ 0.1 mag for all three filters. We have also investigated variations of the completeness limits due to the increased background close to the galaxy centre. A method similar to that described above was adopted, with artificial point sources being added within successive elliptical rings centred on NGC6702 and having the same mean ellipticity and position angle as the galaxy.

| Name      | RA (J2000) | DEC (J2000) | l   | b   | type | heliocentric velocity (km s⁻¹) | distance modulus | B_V (mag) | V_I (mag) |
|-----------|------------|-------------|-----|-----|------|-------------------------------|------------------|-----------|-----------|
| NGC 6702  | 18^h 46^m 57.6^s | +45° 42' 20" | 75.03° | 19.79° | E3 | 4727 | 34.00 | 13.04 | 12.15 |

Table 1. Observational properties of NGC 6702 obtained from NED. The distance modulus is estimated assuming H₀ = 75 km s⁻¹ Mpc⁻¹.

Figure 2. B-band surface brightness profile of NGC 6702. The dashed lines are the 1σ uncertainty envelopes. We define the equivalent radius r_eq = a√1−ε, where a and ε are the semi-major axis and ellipticity respectively.
the population synthesis models of Worthey (1994) we find that colour range of the Galactic GCs. Additionally, using the for globular cluster selection is chosen to span the $B$ occupied by the Galactic GCs (Harris 1996; corrected for cutoffs. Also shown are the regions of the parameter space direction indicated by the arrows in Figures 5 and 6.

$\epsilon = 0.2$, $PA=145^\circ$). The effect of galaxy background was found to be small except for $r_{eq} \lesssim 25''$. However, this is close to the saturated region of the CCD images which is masked out in our analysis. Therefore, the lower sensitivity close to the galaxy center will not affect the results from this study.

**3 SELECTION OF GLOBULAR CLUSTER CANDIDATES**

Globular cluster candidates are selected from the matched catalogue on the basis of their expected range in colour and magnitude. In particular, the full range of possible globular cluster metallicities is $-2.5 < [Fe/H] < +1.0$. Using the colour–metallicity relation of Couture et al. (1990) and Kissler-Patig et al. (1998), the above mentioned metallicity interval corresponds to $1.2 < B - I < 2.5$ and $0.6 < V - I < 1.7$. Figures 3 and 4 plot the $B - I$ and $V - I$ colours against $B - V$ respectively, for the objects in the matched catalogue. The horizontal lines in these figures represent the above mentioned $B - I$ and $V - I$ colour cutoffs. Also shown are the regions of the parameter space occupied by the Galactic GCs (Harris 1996; corrected for Galactic reddening). The accepted range of $B - V$ colour for globular cluster selection is chosen to span the $B - V$ colour range of the Galactic GCs. Additionally, using the population synthesis models of Worthey (1994) we find that the $B - I$, $V - I$ and $B - V$ colours of a stellar population of a given age get redder with increasing metallicity in the direction indicated by the arrows in Figures 3 and 4.

Globular cluster candidates are selected on the basis of their photometric colours if they lie within the regions indicated in Figures 3 and 4, adopting a $3\sigma$ uncertainty on each colour given by the PHOT software. A total of 374 candidate GCs are selected. These objects are then visually inspected to discard resolved sources, likely to be galaxies, and sources lying near hot pixels. The final list consists of 205 candidate GCs. Although visual inspection is not an objective method for identifying point sources, we believe that to the first approximation we successfully eliminate background galaxies from the GC catalogue. This is demonstrated in Figure 5 plotting an estimator of the spatial extent of a source as a function of $I$-band magnitude. The size of a source is estimated by the difference between the $I$-band magnitude measured within 3 and 8 arcsec radius apertures. It is clear from Figure 5 that visually classified resolved sources are separated from point-like objects, particularly at bright magnitudes. At fainter magnitudes ($I > 33$), however, the visual classification scheme becomes more uncertain. Also note that point sources brighter than $I \approx 20$ are likely to have saturated cores.

Additionally, we apply a bright magnitude cutoff set to $\approx 2$ brighter than the most luminous Milky Way GC ($2\text{-Cen}; M_V = -10.29$), shifted to the distance of NGC 6702. This is because NGC 6702 has a young stellar component, with an estimated spectroscopic age of $\approx 2$ Gyrs (Terlevich & Forbes 2000). Therefore, NGC 6702 may possess a number of young metal-rich GCs, expected to be significantly brighter than the old metal-poor GCs. For example, the Worthey (1994) models predict a 2.0 Gyrs old stellar population to be brighter by as much as 1.5 mag in the $I$-band compared to an old metal-poor stellar popula-

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The number counts and colour distribution of foreground stars are estimated using the model developed by Bahcall & Soneira (1980)\footnote{The code for the star count model can be obtained from\url{http://ascl.net/bsgmodel.html}}. In this model the Galaxy consists of both a disk and a spheroid component from which the distribution of stars in the $B$ and $V$-bands as well as the $B - V$ colour distribution are calculated. To predict both the star counts in the $I$-band and the $B - I$ and $V - I$ colour distributions we use the empirical tight relation between the $B - V$ and $V - I$ colours for stars (Bessell 1990). We predict about 120 stars within the field-of-view in the magnitude range $22.5 < B < 26.0$, $21.5 < V < 25.5$ and $20.5 < I < 23.9$ mag. The colour and magnitude selected and visually inspected sample comprises a total of 151 GC candidates. It should be noted that including GCs brighter than the 50% completeness limit for point sources does not change any of our conclusions.

Although we attempt to minimise the fraction of foreground stars and background galaxies in the final GC sample, it is likely that a number of contaminating sources are also selected.

The number counts and colour distribution of foreground stars are estimated using the model developed by Bahcall & Soneira (1980)\footnote{The code for the star count model can be obtained from\url{http://ascl.net/bsgmodel.html}}. In this model the Galaxy consists of both a disk and a spheroid component from which the distribution of stars in the $B$ and $V$-bands as well as the $B - V$ colour distribution are calculated. To predict both the star counts in the $I$-band and the $B - I$ and $V - I$ colour distributions we use the empirical tight relation between the $B - V$ and $V - I$ colours for stars (Bessell 1990). We predict about 120 stars within the field-of-view in the magnitude range $22.5 < B < 26.0$, $21.5 < V < 25.5$ and $20.5 < I < 24.0$ mag. However, the colour selection (assuming a 3σ uncertainty of 0.2 mags) eliminates the majority of stars from the final catalogue. Indeed, the model of Bahcall & Soneira (1980), after both magnitude and colour selection, predicts a total of about ≈ 20 stars in the final GC sample (total of 151 sources). Therefore, the star contamination is about 13%.

To quantify the contamination of the final GC sample by background galaxies we should first estimate the completeness limits for galaxies at different wavebands. Since galaxies are resolved sources, their completeness limit is expected to be at brighter magnitudes to that of point sources (e.g. GCs in NGC 6702). Figure 8 plots the number counts for all the sources detected in our $B$-band frame (total of 1455). Also shown is the star count prediction from the model of Bahcall & Soneira (1980). The difference be-
Figure 7. Size parameter against $I$-band magnitude. The size parameter, estimating the extent of a source, is defined as the $I$-band magnitude difference within 3 and 8 arcsec radius apertures. It is clear that visually classified resolved sources (filled circles) are reasonably well separated from point-sources (open circles), although there is increasing scatter at fainter magnitudes.

Figure 8. Colour–magnitude diagram for the 205 candidate GCs, after colour selection and visual inspection. No magnitude cutoff has been applied. The continuous line corresponds to the completeness limit $I = 23.9$ mag. The dashed lines are the $B$-band magnitude cutoffs for the selection of the final candidate GC sample.

tween the total observed source counts and the predicted star counts in a given magnitude bin estimates, to the first approximation, the number of background galaxies detected in our field (here we ignore the contribution from GCs to the counts). Also shown in Figure 3 are the $B$-band galaxy counts from Metcalfe et al. (1991, 1995). It is clear that the predicted galaxy counts in our field are in good agreement with those from Metcalfe et al. (1991, 1995) to the magnitude $B \approx 25$ and then decline due to incompleteness. We estimate our $B$-band galaxy catalogue to be about $\approx 50\%$ and $\approx 30\%$ complete at $\approx 25.5$ and $\approx 26.0$ mag respectively. Figures 4 and 5 are similar to Figure 3 for the $V$ and $I$-bands respectively. From these Figures, we estimate our $V$- and $I$-band galaxy catalogues to be $\approx 50\%$ complete at $V \approx 24.5$ and $I \approx 23.5$ mag respectively.

Ignoring GCs when calculating the number of galaxies within the field–of–view at $B$, $V$ and $I$–bands, overestimates the galaxy numbers at faint magnitudes. Consequently, the above mentioned completeness limits are conservative estimates.

To quantify the contamination of the GC sample due to background galaxies we use the multiwavelength ($UBVI$) faint galaxy catalogue of Lilly, Cowie & Gardner (1991) complete to $I_{AB} = 24.5$. The AB system is defined as $AB = 48.6 - 2.5 \log f_{\nu}$, where $f_{\nu}$ is the flux cgs units. The transformations from the AB system are: $B \approx B_{AB} + 0.17$, $V \approx V_{AB}$ and $I \approx \chi_{AB} - 0.48$ (Lilly, Cowie & Gardner 1991). Using this galaxy sample we predict a total of $\approx 530$ galaxies with $22.5 \lesssim B \lesssim 26$, $21.5 \lesssim V \lesssim 25.5$ and $20.5 \lesssim I \lesssim 24$ mag after converting to our magnitude system and scaling to the area of the LRIS field–of–view ($\approx 31$ arcmin$^2$). Taking into account the galaxy incompleteness of our sample, estimated from Figures 6 and 7, we predict a total of $\approx 320$ galaxies within the above mentioned magnitude limits. Colour selection, using the aperture colours of Lilly, Cowie & Gardner (1991) and assuming a 3σ uncertainty of 0.2 mags, reduces the number of contaminating background galaxies in our candidate GC sample to $\approx 160$.

However, as demonstrated in Figure 8, visual inspection is efficient in removing extended sources from the sample especially at bright magnitudes. Indeed, a total of $\approx 120$ resolved sources with $22.5 \lesssim B \lesssim 26$, $21.5 \lesssim V \lesssim 25.5$ and $20.5 \lesssim I \lesssim 24.0$ mag have been identified, leaving about $\approx 40$ unidentified galaxies within the final GC sample. This corresponds to a contamination of the GC candidate catalogue (total of 151) by galaxies of $\lesssim 25\%$. We stress that this fraction is an upper limit. In particular, we simulated galaxy bulges following $r^{1/4}$ law profile, with absolute luminosity and effective radius of $M_B = -21.0$ mag and $r_{eff} = 4$ kpc respectively, typical of $L^*$ ellipticals. We found that only half of them are detected to the limit $B = 25.5$ mag and this fraction drops to 10% at $B = 26.0$ mag. Moreover, disk-dominated galaxies will be missed by the detection algorithm at even brighter magnitudes. These surface brightness effects are likely to be significantly less severe in the deep surveys used to determine the galaxy counts at faint magnitudes. Indeed these surveys have effective exposure times as much as 10 times longer than those used here and are therefore, specifically designed to detect sources at faint magnitudes. Since bulges and compact galaxies account for less than 20% of the galaxy population at faint magnitude limits (Elson & Santiago 1996), we anticipate the galaxy
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Figure 9. Differential number counts against $B$-band magnitude for all the sources detected in the NGC 6702 $B$-band frame (histogram). The dashed line is the star counts predicted by the model of Bahcall & Soneira (1980). The open circles are the estimated number of galaxies on the NGC 6702 $B$-band frame (see text). Filled circles are the $B$-band galaxy counts of Metcalfe et al. (1991, 1995). It is clear that at bright magnitudes the estimated number of galaxies detected on the NGC 6702 $B$-band frame are in fair agreement with those predicted by Metcalfe et al. (1991, 1995). At faint magnitudes ($B > 25.0$ mag) incompleteness is affecting the present galaxy sample.

contamination of our final candidate GC sample to be much less than $\approx 25\%$.

4 RESULTS

4.1 Colour distributions

The $B-I$ colour distribution for the final GC sample is plotted in Figure 13. There is evidence for bimodality with the peaks of the distribution at $B-I \approx 1.50$ and $B-I \approx 2.15$. Indeed the KMM statistical test (Ashman et al. 1994) rejects the single Gaussian model at a confidence level better than 99.99%, giving peaks at $B-I = 1.52$ and 2.16 respectively. It should be noted that this result is obtained after clipping objects with $B-I > 2.6$ mag, likely to be background galaxies (3 out of 151; see Figure 13). However, this is not expected to artificially increase the likelihood of the two Gaussian fit. In particular, Ashman et al. (1994) argue that clipping of the data before applying the KMM test is sometimes desirable to remove background contamination. Also shown in Figure 13 is the reddening-corrected colour distribution of Galactic GCs that comprise old metal-poor GCs. The blue peak of NGC 6702 GCs lies close to the Milky Way GC distribution.

Figure 10 plots the $V-I$ colour histogram of both the NGC 6702 GC candidates and the Galactic GCs. There is no strong evidence for bimodality, with the KMM test finding no improvement of the two-group fit over a single Gaussian. This seems surprising since the $B-I$ colour distribution shows clear evidence for bimodality. It is likely that the absence of bimodality in Figure 10 is due to the smaller wave-
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Figure 11. Differential number counts against $I$-band magnitude for all the sources detected in the NGC6702 $I$-band frame (histogram). The dashed line is the star counts predicted by the model of Bahcall & Soneira (1980). The open circles are the estimated number of galaxies on the NGC6702 $I$-band frame (see text). The continuous line is the best fit to the faint $I$-band galaxy counts estimated by Woods & Fahlman (1997). At bright magnitudes the estimated number of galaxies detected on the NGC6702 $I$-band frame are in fair agreement with those predicted by Woods & Fahlman (1997). At faint magnitudes ($I > 22.5$ mag) incompleteness is affecting the galaxy counts from our study.

4.2 Age of the NGC 6702 Globular Clusters

To constrain the age and metallicity of NGC6702 GC system we employ a method similar to that developed by Whitmore et al. (1997) that uses the colour-magnitude relation of GCs. The basic assumptions of the method are discussed in detail by Whitmore et al. (1997). In brief, it is assumed that the GCs comprise two distinct sub-populations of different age and metallicity but with the same initial mass functions. An additional assumption is that both sub-populations have similar numbers of GCs. The older population is taken to have properties similar to those of the Milky Way GCs and therefore consists of old (15 Gyrs) metal-poor ([Fe/H] = −1.5) GCs. The younger population is formed at

Figure 12. $B - I$ colour distribution for the candidate globular clusters, selected on the basis of their colours (see Figures 5 and 6) and magnitudes. Also shown is the colour distribution of Galactic GCs (hatched histograms). The blue peak of the NGC6702 GCs lies close to the Milky Way GC distribution.

Figure 13. $B - I$ histogram for the candidate globular clusters. The filled histogram is the distribution of contaminating stars predicted from the models of Bahcall & Soneira (1980). The hatched histogram is the predicted colour distribution of the contaminating galaxies within the candidate GC sample.
a later stage of the evolution of the galaxy, through either a
merger event (e.g. Ashman & Zepf 1992) or a secondary in
situ star-formation burst (Forbes et al. 1997). The young
GCs, assumed to be metal–rich, are initially bluer than
the old metal–poor ones but become redder after about 1–
2 Gyrs. Similarly, the newly formed GCs are much brighter
than the pre-existing metal–poor population but become
fainter at later stages as they evolve passively. Therefore,
the difference between the colours and magnitudes of the
two GC populations can be used to constrain the age of the
new GCs.

Here we employ the models of Worthey (1994) to predict
the colour and magnitude difference of a young GC popu-
lation relative to an old (15 Gyrs) metal–poor ([Fe/H] = −1.5)
population. The predicted colours of this old stellar popu-
lation are $V - I = 0.921$ and $B - I = 1.564$. The model
predictions for the $B - I$ colour against both $B$ and $I$ mag-
nitudes are shown in Figures 14 and 15 respectively. For
NGC 6702 the difference $\Delta(B - I)$ between the colours of
the two GC populations is taken as the difference be-
tween the peaks of the bimodal $B - I$ distribution in Figure 12.
The blue peak in that figure is assumed to correspond to the
old metal–poor population. This is a reasonable assump-
tion since the blue peak overlaps with the Milky Way GCs fairly
well. The peak colour difference from the KMM statistical
test is $\Delta(B - I) = 0.64 \pm 0.10$, where the error is estimated
assuming an uncertainty of 0.07 mag in defining the peaks
of the bimodal distribution.

The $B$ and $I$-band magnitude difference, $\Delta B$ and $\Delta I$,
between the two GC populations is estimated following
the method of Whitmore et al. (1997) by taking the off-
set between the 10th brightest cluster of the blue and red
subpopulations. The blue and red subpopulations are de-

defined to have colours in the range $1 \leq B - I < 1.8$ and
$1.8 \leq B - I < 3$ respectively. The uncertainties in $\Delta B$ and
$\Delta I$ are estimating using the bootstrap resampling tech-
nique. We find $\Delta B = -0.05 \pm 0.29$ and $\Delta I = -0.84 \pm 0.27$, i.e. the
young GCs are brighter than the old ones.

The results are shown in Figures 14 and 15, and sug-

gest that NGC 6702 GCs have solar to supersolar metallicity
and an age in the range 2-5 Gyrs. Also we note that despite
the errorbars, in both figures the data point lies close to the
age of 3 Gyrs and [Fe/H]=0.5 model prediction. Moreover,
although there are uncertainties in the calculation of the age
of the newly formed GCs, our estimate is in fair agreement
with the galaxy’s spectroscopically estimated age and metal-
licity of $\approx 2$ Gyrs and [Fe/H] $\approx 0.5$ respectively (Terlevich
& Forbes 2000).

To assess the sensitivity of the age and metallicity es-

timates to the adopted single stellar population model of
Worthey, we repeat our calculations using the models de-
veloped by Bruzual & Charlot (1996) and Vazdekis et al.
(1996) with a Salpeter IMF. Both models predict an age
of about 2 Gyrs for the recently formed GC population. The
predicted metallicity of [Fe/H] $> 0.5$ is higher than that from
the Worthey model. Despite, the discrepancies, it is clear
that all three models used here predict that the blue GCs
sub-population is relatively young with a supersolar metal-
licity.

Additionally, assuming an age of 3 Gyrs and [Fe/H]=0.5
for the red GCs, the Worthey (1994) models predict a $V - I$
colour difference between the two GC subpopulations of

\begin{figure}
\centering
\includegraphics[width=\textwidth]{histogram}
\caption{$V - I$ colour distribution for the candidate globular
clusters, selected on the basis of their colours and magnitudes.
Also shown are the colour distributions of Galactic GCs (hatched
histograms).}
\end{figure}

\begin{figure}
\centering
\includegraphics[width=\textwidth]{histogram2}
\caption{$V - I$ histogram for the candidate globular clusters.
The filled histogram is the distribution of contaminating stars pre-
dicted from the models of Bahcall & Soneira (1980). The hatched
histogram is the predicted colour distribution of the contamini-
ging galaxies within the candidate GC sample.}
\end{figure}
Δ(V − I) ≈ 0.3 mag. As discussed in the previous section such a small colour difference is difficult to detect with the present dataset.

4.3 The surface density and specific frequency

We study the surface density profile, ρ(r), of the NGC 6702 GCs using all the sources detected in the B-band to B = 26.1 mag (total of 1237), the 80% completeness limit for point sources. We use the B-band detections instead of the colour selected GC sample because the selection criteria (i.e., completeness limits) are easier to quantify and do not depend on the GC colours. The sample is binned into elliptical rings centred on NGC 6702 and having the same mean ellipticity and position angle as the galaxy (e = 0.2, PA=145°).

The counts in each bin are corrected for geometric incompleteness due to foreground saturated stars, bad columns and the limited size of the field-of-view. Additionally, we mask out the central saturated part of NGC 6702 for radii r_{eq} < 21″. The correction for magnitude incompleteness to the counts in individual bins was found to be small (≈ 3%) and was ignored. The density profile is plotted in Figure 18(a). The background surface density level, ρ_{bg}, was determined by combining the counts within the outermost three bins lying at r_{eq} > 120″ (corresponding to galactocentric distances > 37 kpc) where the density profile reaches a constant level. We find ρ_{bg} = (1.28 ± 0.05) × 10^9 sources deg⁻² (assuming Poisson statistics) or 1114 ± 44 sources over the LRIS field-of-view of 31.2 arcmin² (after applying geometrical corrections). The total number of objects to B = 26.1 mag is 1237 and therefore the number of candidate GC is estimated to be 123±56 (Poisson statistics). In section 3 we found a total of 151 candidate GCs, selected on the basis of their colours and magnitudes in the range 22.5 < B < 26.1 mag. This is in fair agreement with the number of candidate GCs estimated here.

Figure 18(b) shows the candidate GC density profile, estimated by statistically subtracting the constant background surface density level from the density profile of all the sources detected on the B-band frame to B = 26.1 mag. Assuming a power law density profile of the form ρ ∝ r^{-α} for the candidate GCs, we find α = −1.15 ± 0.52. This slope for the GC system of NGC 6702 is consistent with that for other galaxies of a similar magnitude (Forbes et al. 1997). Also shown in Figure 18(b) is the B-band surface brightness profile of NGC 6702. It is clear that the starlight of NGC 6702 drops more steeply than the GC density profile as is often seen in other galaxies (Harris 1991; Durrell et al. 1996). This is often attributed to further dissipative collapse of the host galaxy after the older GCs had formed and before the halo stars formed (e.g., Harris 1986). The radial distributions of the red and blue subpopulations of the GC candidates selected on the basis of their colours and mag-
The globular cluster system of the young elliptical NGC 6702

Figure 18. (a) Left panel: the surface density profile of all the sources detected in the NGC 6702 B-band frame. The continuous lines is the mean background surface density, estimated by combining the counts in the bins at galactocentric distances $>120''$. The dashed lines are the errors around the mean assuming Poisson statistics. (b) Right panel: the surface density profile of the globular cluster candidates (open circles), derived by statistically subtracting the background surface density level from the density profile of the left panel. The best fit power law to the observed profile (continuous line) has the form $\rho \propto r^{-1.15\pm0.52}$. The filled circles are the B-band surface brightness profile of NGC 6702 in arbitrary units. The starlight is more concentrated than the GC system.

magnitudes have also been investigated and have been found to be consistent within the uncertainties.

The globular cluster specific frequency $S_N$ is defined (Harris 1991)

$$S_N = N_T \times 10^{0.4(M_V+15)},$$

where $N_T$ is the total number of globular clusters in the galaxy and $M_V$ is the host galaxy absolute V-band luminosity. To calculate $N_T$ we assume that the B-band luminosity function of NGC 6702 GCs is a Gaussian with the same parameters as that of the Milky Way GCs ($M_{B,\text{peak}} = -6.50\text{mag}$ and $\sigma_B = 1.13$). We note that differences in the peak B-band magnitude of the blue and red GC subpopulation luminosity function are small (see section 4.2). At the distance of NGC 6702 $M_{B,\text{peak}}$ corresponds to $B=27.5\text{mag}$. Previously, we found that the number of GCs to the limiting magnitude $B=26\text{mag}$ is $123\pm56$. Therefore integrating over the entire area of the luminosity function we find $N_T = 1776 \pm 926$. This implies $S_N = 2.3 \pm 1.1$, assuming $M_V = -21.9\text{mag}$ for NGC 6702.

It is an interesting exercise to predict the evolution of $S_N$ with time to investigate whether NGC 6702 will resemble normal ellipticals within a few Gyrs. For this calculation we assume that the total number of GCs is conserved and that the change in $S_N$ is due to passive evolution of the host galaxy. Furthermore, we assume that during the most recent major star-formation event, that took place about 2 Gyrs ago, 10% of the NGC 6702 mass (see section 5 below) has been converted into metal-rich ([Fe/H]=+0.5) stars. After 10 Gyrs this new stellar population will fade in the V-band by $\Delta V = 0.18\text{mag}$ (Worthey 1994), corresponding to $S_N = 2.7$. Harris, Harris & McLaughlin (1998) found a mean of $S_N = 5.4 \pm 0.3$ for large ellipticals in the Virgo, Fornax and the NGC 5846 group (excluding cD galaxies). However, exclusion of Bright Cluster Galaxies (BCGs) results in a mean of $S_N = 3.3 \pm 0.2$, in better agreement, although somewhat larger, with the estimated $S_N \approx 2.7$ for NGC 6702 after 10 Gyrs. However, the estimated $S_N$ evolution is sensitive to the galaxy mass fraction involved in the most recent major star-formation event, which in turn is uncertain and difficult to constraint. We calculate an upper limit for the NGC 6702 $S_N$ after 10 Gyrs by assuming the extreme scenario where as much as 50% of the galaxy mass has been turned into stars. We estimate a fading of $\Delta V = 0.82\text{mag}$

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Table 2. Comparison between the properties of NGC 6702 and NGC 1700 ($H_0 = 75 \text{ km s}^{-1} \text{ Mpc}^{-1}$).

| Type     | NGC 6702 | NGC 1700 |
|----------|----------|----------|
| Mg$_V$ (mag) | -21.9    | -21.5    |
| Environment | Field   | Group    |
| $V/\sigma^+$ | 0.21     | 0.38     |
| Spectroscopic age (Gyr) | 2        | 2.3      |
| [Fe/H] 2    | 0.5      | >0.5     |
| [Mg/Fe] 2   | 0.12     | 0.13     |
| Fundamental Plane residual | -0.20    | -0.37    |
| $S_N^2$     | 2.3 ± 1.1| 1.4$^{+1.0}_{-0.6}$ |

1Bender et al. 1992
2Terlevich & Forbes 2000
3Forbes, Ponman & Brown 1998
4The $S_N$ for NGC 1700 from Brown et al. 2000

(Worthey 1994) within 10 Gyrs, corresponding to an upper limit of $S_N = 4.9$. Therefore, we conclude that the NGC 6702 $S_N$, although lower than that of bright cluster ellipticals, is likely to be similar to that of ‘normal’ (old) present-day ellipticals (excluding BCGs) within the next 10 Gyrs. The calculations above are also consistent with suggestions that galaxy interactions and mergers although unable to account for large $S_N$ galaxies, may explain the low values of $S_N \approx 2$ in relatively isolated ellipticals (Elmegreen 1999).

5 DISCUSSION

A comparison between the properties of NGC 6702 and the group proto-elliptical NGC 1700 (Brown et al. 2000; Whitmore et al. 1997), is shown in Table 2. It is clear that there are quite a few similarities between the two galaxies. Both of them have similar optical luminosities and deviate from the Fundamental Plane in the sense of having younger ages on average (Forbes, Ponman & Brown 1998). This is confirmed by their central spectroscopic ages of $\sim 2$ Gyrs (Terlevich & Forbes 2000). The central stars appear to be of high metallicity ([Fe/H] $\sim 0.5$) with supersolar [Mg/Fe] abundance. Supersolar abundances have been the source of much debate in the literature but probably indicate either a rapid star formation time-scale or an IMF skewed towards massive stars. Both galaxies have radial velocity to velocity dispersion ratios, $(V/\sigma)^+$, indicating that they are anisotropic rotators. Bender et al. (1992) have suggested that the last major merger for anisotropic galaxies involved relatively little gas, i.e. the progenitors were largely stellar. In the case of NGC 6702 and NGC 1700 there must have been sufficient gas to support the central starburst and associated cluster formation. The similarities above also suggest a similar formation mechanism for these galaxies.

Moreover, our analysis above indicates the presence of two GC populations in NGC 6702 with different metallicities and ages. It is an interesting exercise to estimate the formation efficiency of the young (red; $1.8 \leq B - I \leq 3$) GC subpopulation, assuming they formed during the most recent star-formation burst, and compare it to that of stars formed in the same burst.

A robust estimator of the GC formation efficiency is the ratio of the total GC mass to the total baryonic galaxy mass. For this calculation we assume a typical GC mass of $3 \times 10^6 M_\odot$ (Harris, Harris & McLaughlin 1998) and a mass-to-light ratio for NGC 6702 $(M/L)_V = 8 M_\odot/L_\odot$, typical to that of ellipticals (Faber & Gallagher 1979; Binney & Tremaine 1987). We have identified a total of 88 red GCs $(1.8 \leq B - I \leq 3.0)$ to the limit $B \approx 26.0$ mag, assumed to be the newly formed metal-rich GCs of NGC 6702. We adopt a Gaussian luminosity function for this GC subpopulation with parameters similar to those used in section 4.4 ($(M_B,peak) = -6.50$ mag and $\sigma_B = 1.13$). This is a reasonable assumption, since the blue and red GC subpopulations of NGC 6702 in section 4.4 have a similar $B$-band magnitude distribution within few tenths of a magnitude. Integrating over the entire luminosity function, assuming a completeness limit for our sample $B \approx 26.0$ mag, we estimate a total of 970 red GCs. Thus, the total mass of the red GCs is $M_{GC} = 970 \times 3 \times 10^5 M_\odot = 2.9 \times 10^9 M_\odot$. The total baryonic mass in stars of NGC 6702, estimated using its V-band luminosity and the assumed mass-to-light ratio, is found to be $M_V = 3.6 \times 10^{11} M_\odot$. Therefore, we estimate that 0.08% of the NGC 6702 mass has been turned into GCs during the most recent star-formation event, which is lower compared to recent estimates for the GC formation efficiency of 0.25% (Harris, Harris & McLaughlin 1998; McLaughlin 1999). It should be noted that this calculation assumes that the estimated baryonic mass in stars represents the total mass of the system, i.e. most of initial gas has been used for star-formation.

Elmegreen et al. (1999), argue that GCs represent about 1% of the total galaxy star-formation. Therefore, for the GC formation efficiency of 0.08% we expect that $\approx 8\%$ of the total NGC 6702 mass has turned into field stars (not in bound clusters) during the most recent star-formation event. Is our estimate of the mass in new stars (i.e. 8% of the total) consistent with the spectral line indices in the central $r_\ast/8$ region (Gonzalez et al. 1993)? If we assume that the new stars formed with a metallicity of [Fe/H]=$-0.5$, and the old stars (representing 92% of the central stellar mass) have solar metallicity and an age of 15 Gyrs, then the new stars would have to be slightly younger than 1 Gyr old. Assuming a time since the starburst of $\approx 2$ Gyrs and applying that to the new stars, their mass fraction would have to be increased to $\approx 15\%$. Furthermore, if the galaxy reveals evidence for a young stellar population outside of the central region this will also point towards a large mass fraction (see also Trager 1997). Clearly there are large uncertainties involved in the above arguments, but there is some agreement between the mass fraction of young stars estimated independently using spectroscopic line indices and the properties of the NGC 6702 GC system. Moreover, it seems clear that a large fraction of new stars formed in NGC 6702 a few Gyrs ago. In the merger context, it implies a high gas fraction for the progenitors and hence a highly dissipational event. This is contrary to the claims of Bender et al. (1992) that anisotropic galaxies (such as NGC 6702 and NGC 1700) formed by largely dissipationless mergers.

Next, we compare the fraction of galaxy starlight in newly formed dense clusters relative to that of young field
The globular cluster system of the young elliptical NGC 6702

6 CONCLUSIONS

We have carried out B, V and I-band photometric observations of the elliptical galaxy NGC 6702 to explore its GC system. The GC candidates are selected on the basis of their colour, magnitude and morphology. We find strong evidence that the B − I colour distribution of GC candidates is bimodal with the blue peak having a colour similar to that of the old metal-poor Galactic GCs. Assuming that the blue GC subpopulation is indeed old and metal-poor, we estimate that the red GCs have an age of 2–5 Gyrs and supersolar metallicity. Despite the large uncertainties, this is in reasonable agreement with the spectroscopically estimated age and metallicity of about 2 Gyrs and 0.5 respectively for the galaxy. Thus the red GCs in NGC 6702 are consistent with having formed during the most recent starburst, likely to have been triggered by a merger or an accretion event, a few Gyrs earlier. Additionally, we find a specific frequency for the NGC 6702 GC system of SN = 2.3 ± 1.1. We predict that passive evolution of the NGC 6702 stellar population formed during the most recent star-formation event will increase the NGC 6702 specific frequency to ≈ 2.7 within 10 Gyrs, in reasonable agreement with the mean for present-day ellipticals, excluding the largest cluster ellipticals. Finally, using the NGC 6702 GC properties and making reasonable assumptions about the GC formation efficiency, we estimate the mass fraction of newly formed stars during the most recent star-formation event. We find reasonable agreement with the mass fraction estimated using stellar population synthesis models to interpret the line indices of the underlying galaxy starlight.

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