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NUSTAR AND XMM-NEWTON OBSERVATIONS OF 1E1743.1-2843: INDICATIONS OF A NEUTRON STAR LMXB NATURE OF THE COMPACT OBJECT

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ABSTRACT

We report on the results of NuSTAR and XMM-Newton observations of the persistent X-ray source 1E1743.1-2843, located in the Galactic Center region. The source was observed between 2012 September and October by NuSTAR and XMM-Newton, providing almost simultaneous observations in the hard and soft X-ray bands. The high X-ray luminosity points to the presence of an accreting compact object. We analyze the possibilities of this accreting compact object being either a neutron star (NS) or a black hole, and conclude that the joint XMM-Newton and NuSTAR spectrum from 0.3 to 40 keV fits a blackbody spectrum with $kT \sim 1.8$ keV emitted from a hot spot or an equatorial strip on an NS surface. This spectrum is thermally Comptonized by electrons with $kT \sim 4.6$ keV. Accepting this NS hypothesis, we probe the low-mass X-ray binary (LMXB) or high-mass X-ray binary (HMXB) nature of the source. While the lack of Type-I bursts can be explained in the LMXB scenario, the absence of pulsations in the 2 mHz–49 Hz frequency range, the lack of eclipses and of an IR companion, and the lack of a $K_s$ line from neutral or moderately ionized iron strongly disfavor interpreting this source as a HMXB. We therefore conclude that 1E1743.1-2843 is most likely an NS-LMXB located beyond the Galactic Center. There is weak statistical evidence for a soft X-ray excess which may indicate thermal emission from an accretion disk. However, the disk normalization remains unconstrained due to the high hydrogen column density ($N_H \sim 1.6 \times 10^{23}$ cm$^{-2}$).

Key words: accretion, accretion disks – stars: neutron – X-rays: binaries – X-rays: individual (1E1743.1-2843)

1. INTRODUCTION

The X-ray source 1E1743.1-2843 was discovered during the first soft X-ray imaging observation of the Galactic Center, which was performed by the Einstein Observatory (Watson et al. 1981), and has been detected in all of the subsequent observations performed by X-ray satellites with imaging capabilities above 2 keV (Kawai et al. 1988; Sunyaev et al. 1991; Pavlinsky et al. 1994; Lu et al. 1996; Cremonesi et al. 1999; Porquet et al. 2003; Muno et al. 2009; Bird et al. 2010). Its position has been determined with Chandra to be $\alpha_{2000} = 17^h 46^m 21^s 09$; $\delta_{2000} = -28^\circ 43^\prime 42^\prime\prime 67$ with a reported 0.21 sigma accuracy (Evans et al. 2010). Because of its high column absorption ($N_H = 1.3 \pm 0.1 \times 10^{23}$ cm$^{-2}$; see Cremonesi et al. 1999), the source is likely in the Galactic Center ($d = 7.9 \pm 0.3$ kpc; McNamara et al. 2000) or beyond, while the orbital inclination is smaller than 70° (Cowley et al. 1983). The analysis performed by Porquet et al. (2003) detected no pulsations or quasi-periodic oscillations in the 2.4 mHz–2.5 Hz frequency range using EPIC-MOS and the PN fullframe mode (time resolution 2.6 s and 200 ms, respectively). However, since many X-Ray binaries present quasi-periodic variations above 2.5 Hz, the XMM-Newton data were not suitable to probe the millisecond pulsations that could indicate a low-mass X-ray binary (LMXB) nature. In their analysis, Porquet tested several single-component spectral models (i.e., absorbed power law, absorbed blackbody, absorbed disk blackbody) but could not distinguish between these models due to the narrow bandpass from 2 to 10 keV. In this regard, NuSTAR’s ability to perform high-resolution, broadband spectroscopy allows us to probe more deeply the nature of the high-energy emission from this source.

The presence of an accreting object is required to explain the unabsorbed source luminosity, which is of the order of $L_X \sim 10^{37}$ $d_{10}$ erg s$^{-1}$, where $d_{10}$ is the source distance expressed in units of 10 kpc, while the absence of periodic oscillations and eclipses favors a scenario in which a compact object (either a neutron star (NS) or a black hole) accretes matter from a low-mass companion (LMXB systems). LMXBs in this luminosity range which contain an NS are usually characterized by thermonuclear flashes of accreted matter that ignites on the NS surface (Type I X-ray bursts), but these bursts have never been observed for 1E1743.1-2843 in 20 years of X-ray observations. So far, this has led to the conclusion that (1) the accretion rate is high enough to allow stable burning of the accreted material, which would imply an $M$ value comparable to the Eddington limit (Bildsten 2000), and thus a distance greater than 8 kpc, (2) the bursts are suppressed by the presence of intense magnetic fields (at least
10^9 G), and (3) the accreting compact object is a black hole (Porquet et al., 2003; Del Santo et al., 2006). The presence of strong magnetic fields (>10^12 G), however, should be accompanied by cyclotron absorption features and pulsations, neither of which have been observed in 1E1743.1-2843. The source showed marginal variability on month timescales in the 1.3–10 keV energy range (Cremonesi et al., 1999), and less than 18% of variability between 10^{-4} and 2.5 Hz in the 2–10 keV energy range (Porquet et al., 2003).

Here, we present the results of four observations of 1E1743.1-2843 performed with the NuSTAR and XMM-Newton satellites in 2012 September–October. The two satellites provided almost simultaneous broadband X-ray spectroscopy from 0.1 to 79 keV and, thanks to NuSTAR’s timing capabilities, we can also probe pulsations down to 2 millisecond timescales, providing unprecedented opportunities to investigate this source. In Section 2, we describe the observations, and in Section 3 we discuss the timing and spectral analysis. Our results and discussion are presented in Section 4.

2. OBSERVATIONS

All of the observations were performed between 2012 September and October. NuSTAR observed 1E1743.1-2843 during the so-called “mini-survey” of the Galactic Center, which took place shortly after the NuSTAR in-orbit checkout. The observation was performed on 2012 October 15th, and the total exposure time was 26 ks. The source was ~7' off-axis in the NuSTAR observation and appears distorted due to the asymmetric point-spread function (PSF) shape at a large off-axis angle (Madsen et al., 2015), as can be seen in Figure 1.

NuSTAR is the first X-ray satellite with multilayer hard X-ray optics and is operational in the energy range 3–79 keV (Harrison et al., 2013). The mission carries two identical telescopes with grazing incidence optics, each one focusing on separate detector modules, Focal Plane Modules A and B (FPMA, FPMB), at a distance of 10 m. These CdZnTe detectors have a total field of view (FOV) of 13' x 13' (Harrison et al., 2013). The telescope PSF has an 18” FWHM with extended tails resulting in a half-power-diameter of 58" (Harrison et al., 2013).

To improve the low-energy sampling of the source spectrum, we looked for other high-energy observations of 1E1743.1-2843 that were performed during approximately the same period. We adopted three XMM-Newton observations performed in imaging mode during 2012 September for a total exposure time of 90 ks. Observation 0694640401, however, is strongly contaminated by solar emission, and we therefore decided to not use it, reducing the useful exposure time to 65 ks. The XMM-Newton EPIC-pn camera provides data with nominal accuracy in the 0.3–10 keV energy band (XMM-Newton Science Operations Centre Team, 2014), providing a good overlap with the NuSTAR data, and thereby minimizing possible bias in the spectral modeling.
The 
Newton observations are not simultaneous with XMM-
Newton. However, due to the small difference in time (a few
weeks, see Table 1), and since 1E1743.1-2843 does not usually
exhibit substantial spectral variability (Cremonesi et al. 1999),
we jointly fit the two data sets with a cross-normalization factor
to account for any flux variation. During the observations,
the difference in source flux measured by the two instruments in
the overlapping energy band (3–10 keV) was below 1.3%,
indicating a good compatibility of the data sets.

3. DATA ANALYSIS

3.1. NuSTAR

We analyzed the NuSTAR data set (obsID 40010005001)
using the NuSTAR Data Analysis Software (NuSTARDAS)
version 1.3.1 (2013 December 9), HEASOFT 6.15.1, and the
most up-to-date calibration files and responses. The software
applies offset correction factors to the energy response to
account for the movement of the mast which causes a varying
position of the focal spots on the detector planes. For the data
from each of the two modules, the pipeline produces an image,
spectrum, and deadtime-corrected light curve. For each
NuSTAR observation, the source and background subtraction
regions must be carefully evaluated due to the possible presence
of contaminating sources outside the FOV that induce
stray light patterns on the detectors. Unfortunately the NuSTAR
detectors are not entirely shielded from unfocused X-rays, and
this stray light can be significant if there are bright X-ray
sources within ∼2°–3° of the pointing direction (Krivonos et al.
2014; Wik et al. 2014; Mori et al. 2015). In this observation,
FPMB is highly contaminated by two different stray light
patterns, as can be seen in Figure 1, and furthermore the source
focal spot straddles the two different patterns. Due to the
complexity of separating the source emission from the stray
light, we decided to discard all of the data from FPMB.

The source spectra are shown in Figure 2 and were obtained
by extracting photons in an elliptical region of 95° × 46° semi
axis, rotated by 35° clockwise relative to north and centered
on the source centroid, and by subtracting the count rate measured
in a nearby circular background region of radius 114° (see
Figure 1). The different area normalizations were taken into
account in the background subtraction. The dead layer
thickness of the two modules varies depending on the location
on the detector and, since the pipeline could fail to correct the
response matrix for this effect to a suitable accuracy for high
off-axis angles, we decided to ignore the data below 5 keV.
Also, due to imperfect background subtraction, the data above
40 keV were ignored. All of the data bins were grouped to
reach at least 30 counts. The total source count rate in this
energy range is 2.10 ± 0.01 cts s⁻¹.

3.2. XMM-Newton

We extracted the spectra of the EPIC-pn camera in the
0.3–10 keV energy range following the standard procedure
described in the XMM-Newton software analysis guide13 for
observations 0694640401, 0694640501, and 0694641201,
using the SAS software release xmmssas_20131209_1901-
13.5.0. The light curves and spectra of 1E1743.1-2843 were
extracted from annular extraction regions excluding those
zones where the number of counts exceeded 800 to avoid
pileup if present (see Figure 3). The light curves are shown in
Figure 4.

EPIC-pn data below 0.3 keV are mostly related to artifacts
and noise and were excluded (XMM-Newton Science Opera-
tions Centre Team 2016), and all the bins were grouped to
reach at least 30 counts. The mean count rate for the two
observations is 2.62 ± 0.01 cts s⁻¹.

3.3. Timing Analysis

A light curve of the NuSTAR observation with 10 ms time
bin resolution was extracted for FPMA in the energy14 range
3–60 keV using nuproducts. Timing data were corrected for
deadtime and for arrival times at the Solar System barycenter
using the JPL 2000 ephemeris (Standish 1982; for this purpose,
we used the barycorr tool in the HEASOFT 6.16 distribution).
We then calculated the power spectra on different contiguous
time intervals and averaged them into a total spectrum. Each
single spectrum was built using intervals of 32,768 bins, and
the total spectrum was built averaging 87 intervals in a single
frame. This was finally rebinned in frequency channels for
more statistics. An offset constant term was subtracted from the
total spectrum to remove the Poisson noise level and
compensate for the effects of the deadtime correction.

We also performed a Lomb-Scargle periodogram of the
same frequency interval as the power spectrum and found no
significant signal. The power spectrum and Lomb period
analysis were also performed for the two XMM-Newton
observations, also yielding no positive detection of
periodicities.

13 ftp://legacy.gsfc.nasa.gov/xmm/doc/xmm_abc_guide.pdf The XMM-
Newton ABC Guide: An Introduction to XMM-Newton Data Analysis.
14 The low-energy range is different from that used in the spectral analysis
section, in the attempt to maximize the S/N.
3.4. Spectral Analysis

We analyzed the NuSTAR and XMM-Newton data sets using XSPEC (Arnaud 1996) version 12.8.0, using the TBabs with abundances set as in Wilms et al. (2000), and cross-sections set as in Verner et al. (1996) to model the effect of X-ray absorption. We performed the fit allowing for the normalizations among the three different observations to vary. These normalizations relative to the different XMM-Newton data sets remained constant for every model and were found to be $C_1 = 0.99 \pm 0.01$, $C_2 = 1.00 \pm 0.01$ for observations 0694640501 and 0694641201, respectively.

We used the following four models to determine the nature of the source, the first three of which were used to test the origin of the low-energy emission.

1. Model 1: a blackbody (bbbodyrad) and a power law with a high-energy cutoff (power law $\times$ highEcut). This is a typical single-component spectral model and assumes that the source is an NS binary.

2. Model 2: a disk blackbody (diskbb) plus a power law $\times$ highEcut. This is also a typical single-component spectral model but assumes that the compact source is a black hole binary.

3. Model 3: a disk blackbody (diskbb), a blackbody (bbbodyrad), and the power law $\times$ highEcut. This model probes deeper into the NS hypothesis, adding an accretion disk component to model 1.

Figure 3. Images of 1E1743.1-2843 from the XMM-Newton EPIC-pn camera. White ellipses show the annular extraction region and green circles show the background regions. The bright central zone of the source was excluded due to pileup in the central and rightmost panels, while the pileup threshold was not exceeded in the leftmost one (thus no inner ellipse was plotted on the image).

Figure 4. Light curves of 1E1743.1-2843 acquired from the NuSTAR FPMA in the 3–80 keV energy range (left, 300 s bins), and from the two XMM-Newton analyzed observations in the 0.3–10 keV energy range (right, 600 s bins, observation 0694640501 and 0694641201 in black and gray, respectively).

Figure 5. Power spectrum of the NuSTAR observation of 1E1743.1-2843, with indication of a low frequency break at frequencies $>0.02$ Hz. Upper limits are $1\sigma$. 
| XSPEC Model | Parameter | Units | Model 1 | Model 2 | Model 3 | Model 4 |
|-------------|-----------|-------|---------|---------|---------|---------|
| TBabs       | $N_H$    | 10$^{22}$ cm$^{-2}$ | $14.8^{+0.9}_{-0.8}$ | $16.6^{+0.5}_{-0.5}$ | $16.5^{+0.3}_{-0.3}$ | $15.3^{+0.8}_{-1.0}$ |
| bbodyrad    | $kT^*$  | keV   | $1.8^{+0.1}_{-0.1}$ | $...$ | $1.8^{+0.1}_{-0.1}$ | $1.8^{+0.0}_{-0.0}$ |
| bbodyrad    | norm$^a$ | ...   | $1.3^{+0.5}_{-0.3}$ | $...$ | $1.0^{+0.3}_{-0.3}$ | $1.7^{+0.4}_{-0.4}$ |
| diskbb      | $kT_{in}$ | keV  | $2.4^{+0.2}_{-0.3}$ | $0.12^{+0.04}_{-0.02}$ | $0.12^{+0.04}_{-0.03}$ | $...$ |
| diskbb      | norm$^a$ | ...   | $1.0^{+0.1}_{-0.0}$ | $0.3^{+0.0}_{-0.0}$ | $1.3^{+0.2}_{-0.2}$ | $...$ |
| power law   | $\Gamma$| ...   | $1.0^{+0.3}_{-0.3}$ | $0.9^{+0.3}_{-0.3}$ | $1.3^{+0.3}_{-0.3}$ | $...$ |
| power law   | norm$^a$ | ...   | $0.009^{+0.01}_{-0.009}$ | $0.01^{+0.01}_{-0.005}$ | $0.02^{+0.01}_{-0.01}$ | $...$ |
| highecut    | cutoffE | keV   | $6.6^{+0.4}_{-0.3}$ | $6.9^{+0.3}_{-0.3}$ | $7.2^{+0.2}_{-0.2}$ | $...$ |
| highecut    | foldE   | keV   | $8.1^{+1.3}_{-1.3}$ | $6.3^{+0.5}_{-0.5}$ | $9.0^{+1.3}_{-1.3}$ | $...$ |
| compTT      | $T_0$   | keV   | ...       | $...$ | $...$ | $0.014^{+0.01}_{-0.02}$ |
| compTT      | $kT$    | keV   | $...$     | $...$ | $...$ | $4.6^{+0.6}_{-0.4}$ |
| compTT      | $\tau_p$| ...   | $...$     | $...$ | $...$ | $6.2^{+4.5}_{-1.4}$ |
| compTT      | norm$^a$| ...   | $...$     | $...$ | $...$ | $0.03^{+0.03}_{-0.01}$ |
| $\chi^2$   | dof$^b$ | ...   | $2793.8$ | $2835.4$ | $2766.4$ | $2785.0$ |
| reduced $\chi^2$ | ...     | $1.061$ | $1.075$ | $1.050$ | $1.057$ | $...$ |

Notes. The errors are expressed with 90% confidence. Parameters not reported were not constrained by the fit.

$^a$ $R_{in}/D_{in}$, where $R_{in}$ is the inner radius in km and $D_{in}$ is the distance to the source in units of 10 kpc.

$^b$ ($R_{in}/D_{in})^2 \cos \theta$, where $R_{in}$ is the apparent disk radius and $\theta$ is the angle of the disk ($\theta = 0$ is face-on).

$^c$ The power-law component normalization in units of photons keV$^{-1}$ cm$^{-2}$ s$^{-1}$.

$^d$ The thermal Comptonization component normalization in units of photons keV$^{-1}$ cm$^{-2}$ s$^{-1}$.

4. Model 4: a disk blackbody (diskbb), a blackbody (bbodyrad), and the compTT Comptonization model.

In this model, the power law in model 3 is replaced by a more physical model.

To model the emission from the accretion disk, we used the multi-color diskbb model mentioned before (Mitsuda et al. 1984; Makishima et al. 1986).

The fit results shown in Table 2 indicate that the source emission is mainly contributed by a prominent blackbody component at ~2 keV plus a high-energy continuum which can be described equally well by an empirical law (power law $\times$ highEcut) or by a thermal Comptonization (compTT) component. As can be seen from Table 2, all of the models provide a good fit for the data. In the following, we will assume model 3 as our baseline, since it provides the lowest $\chi^2$ value, and a reasonable physical interpretation of the data, as discussed in Section 4.

Adding another absorption component with partial covering (pcfabs) to model 1 did not improve the fit results, and therefore we conclude that the source is not partially covered. Also, the spectroscopic data do not require the addition of further high-energy components such as, for instance, reflection, as we verified by adding the comprefl (Ballantyne et al. 2012) reflection component to model 3.

We then addressed the high-energy component taking as a baseline model 3 and replacing the power-law plus cutoff component with a more physical model, specifically, compTT (Titarchuk 1994), which describes the Comptonization of soft photons in a thermal plasma of high-energy electrons above the accreting source. In the resulting model 4, the compTT component gives a worse fit than the power law with an exponential cutoff (model 3). However, model 4 still yields a better fit than models 1 and 2, and has the advantage of allowing a more physical interpretation of the source spectrum than the empirical model 3. Furthermore, we tested if a single thermal component and a Comptonized component could fit the data, and replaced the power law in model 1 with the compTT model. However, this resulted in a worse fit ($\chi^2_{red} \sim 1.07$). Finally, if we test model 4 for the presence of the fluorescence line of neutral iron, $\chi^2_{red}$ rises to 1.2. Therefore, the presence of the iron $K_{\alpha}$ line is also not required. From model 4, we derive an upper limit on the iron $K_{\alpha}$ equivalent width (EW) of 4.9 eV.

Summarizing, model 3 and model 4 give the best results. Model 3 has the lower reduced $\chi^2$ value between the two, although the compTT model allows the determination of the physical parameters of the source. The resulting spectra for model 4 are shown in Figure 6, and the results of the overall fits are reported in Table 2.

4. DISCUSSION

In this work, we have reported on a broadband (0.3–40 keV) spectral analysis of 1E1743.1-2843, which was observed with NuSTAR and XMM-Newton. A similar analysis was previously performed using INTEGRAL data (Del Santo et al. 2006), but the improved energy resolution and sensitivity of NuSTAR allow us to better constrain the hard X-ray continuum, identifying the presence of Comptonization and of a cutoff in the high-energy emission for the first time. This allowed us to use more sophisticated models compared to previous works (Cremonesi et al. 1999; Porquet et al. 2003; Del Santo et al. 2006). Below, we summarize the results obtained and discuss those which support either the LMXB or high-mass X-ray binary (HMXB) nature of the system.

Regarding the low-energy emission of the source, the presence of the $K_{\alpha}$ line is a strong indication of an NS nature of the compact object. The blackbody radius is $\sim 1$ km if we assume the source to be located in the Galactic Center, at $d \sim 8.8$ kpc. This is not compatible with emission from a boundary layer near the NS surface ($R_{BL} - R_{NS} \sim 2$ km for low $M$, Popham & Sunyaev 2001), but is consistent with emission from a restricted region of the NS surface (i.e., from an equatorial belt in the orbital plane, or with magnetically

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**Table 2**

Fit Results of the NuSTAR and XMM-Newton Observations
driven accretion onto polar caps). In the latter case, the magnetic field strength of the compact source should be \(>10^9\) G. There may be soft excess which could be interpreted as weak emission from an accretion disk \((kT \sim 0.1\,\text{keV})\). However, the XMM-Newton data were unable to constrain the blackbody normalization due to the high hydrogen column absorption \((N_H \sim 1.6 \times 10^{23}\,\text{cm}^{-2})\). In the HMXB scenario, this soft excess could also be interpreted as a blend of emission lines, as thermal emission from the NS surface, or Thomson scattering of the hot spot emission by the accreting material (van der Meer et al. 2004). The reliability of the soft excess detection is worth discussing because the \(\chi^2\) improvement it provides is moderate. The use of the F-test could result in unreliable results (Protassov & van Dyk 2002), and so we performed a series of Monte Carlo simulations with the simfext routine to confirm the presence of such a component, following the approach described by Bhalerao et al. (2015). The highest \(\Delta \chi^2\) obtained in 10,000 simulations is 18.7, which is significantly lower than the \(\Delta \chi^2 = 43.2\) obtained by the real data (see Figure 7). We estimate that \(>10^9\) simulations would be required to get \(\Delta \chi^2 > 40\), corresponding to a significance higher than 6\(\sigma\). Nevertheless, even if the presence of the soft excess is statistically preferred, then we cannot exclude that the low-energy spectrum is affected by systematics, which could arise, e.g., in adding two exposures taken at different epochs. We regard this as a result that needs to be confirmed by better-quality, low-energy data from future observations of this source.

The presence of a neutral iron \(K_{\alpha}\) line at 6.4 keV is not statistically required, with an upper limit on the EW of 4.9 eV from model 4 (6.7 eV from model 3). This value is compatible with the expected properties of LMXBs (Asai et al. 2000) and radio-quiet quasars (George et al. 2000), as already pointed out by Porquet et al. (2003). This weakens the HMXB hypothesis for this source, since the iron line is usually detected in such systems. On the other hand, the lack of the Fe \(K_{\alpha}\) line in LMXBs is very common, and is usually associated with the Baldwin effect: the high luminosity of the X-ray source increases the degree of iron ionization (Torrejón et al. 2010).

The high-energy emission is characterized by a power law \((\Gamma \sim 1.3)\) with a high-energy cutoff \((E_{\text{cut}} \sim 7.2\,\text{keV})\), identified for the first time thanks to the NuSTAR hard X-ray sensitivity. The NuSTAR data also unambiguously identified the presence of a Comptonization component induced by an electron population with a temperature of a few keV \((kT_0 \sim 4.6\,\text{keV})\), in accordance with the exponential cutoff value found, indicating a large viewing angle (see Matsuoka & Asai 2013). Even if X-rays are produced near the NS at energies of several tens of keV, the observed X-rays will be shifted to lower energy due to Comptonization by the more distant low-energy plasma with a temperature of several keV. This Comptonization component has a strong interplay with the 1.8 keV blackbody up to \(\sim 20\,\text{keV}\), since the total hard X-ray spectrum cannot be described as a single power law. We tested for the presence of a reflection component, and found that it is not required to explain the data. A fit with a power-law model

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**Figure 6.** Left—folded spectra, backgrounds, and residuals with respect to model 4 for all of the spectra analyzed. Right—the analyzed spectra, unfolded through model 4. The plots are rebinned in groups of 30 channels for display purposes.

**Figure 7.** Results of Monte Carlo simulations for testing the presence of the disk component. We simulated the spectra from model 3 without (our null-hypothesis) and with the disk component. The histogram shows the \(\Delta \chi^2\) obtained in 10,000 simulations (in black), together with the \(\Delta \chi^2\) distribution expected from the addition of 2 free parameters (red line). The vertical dashed line is the \(\Delta \chi^2\) value obtained from the actual data (43.19).
without breaks brought higher values of the photon index ($\Gamma = 2.3^{+0.09}_{-0.10}$), which are compatible with previous results (Cremonesi et al. 1999; Porquet et al. 2003; Del Santo et al. 2006), but also produced a higher value of the $\chi^2$/dof $= 2880.79/2639 = 1.09$ compared to that obtained with the high-energy cutoff ($\chi^2_{\text{red}} = 1.05$; see Table 2), indicating that using the former description for the high-energy data is marginally justified at best.

The source luminosity is $L_{2-10 \text{ keV}} \sim 1.5 \times 10^{36} d_{10 \text{kpc}}^2 \text{ erg s}^{-1}$, so that if we assume that the source is located in the Galactic Center, then we obtain $L_{2-10 \text{ keV}} \sim 10^{36} \text{ erg s}^{-1}$, which is within the typical range of luminosity for X-ray bursters (Cremonesi et al. 1999; Bildsten 2000). To investigate further, we have extracted the $K$-band image of the Vista survey “VISTA Variables in the Via Lactea (VVV)” (Minniti et al. 2010; see Figure 8). Both the XMM-Newton and Chandra positions of 1E1743.1-2843 are visible in the image, as is the position of the source UGPS J174621.12284343.3 as reported in the UKIDSS catalog (coordinates: R.A. = 17:46:21.13, $\delta = -28:43:43.17$ J2000). The Chandra error circle excludes the possibility of this source being the counterpart.

In the LMXB hypothesis, the fact that not a single burst has been observed in over 20 years suggests that the source is a rare burster (in’t Zand et al. 2004). In principle, the lack of bursts could be explained if the source is located outside the Galaxy. Specifically, to reach $L = 2.9 \times 10^{37} \text{ erg s}^{-1}$, corresponding to the stable burning of accreted material, 1E1743.1-2843 would need to be placed at $d \sim 40 \text{ kpc}$. However, a study of the stability conditions for accreting objects (Narayan & Heyl 2002) indicates that, for an NS with a surface temperature of $\sim 2 \text{ keV}$, there is a luminosity range where stable accretion is possible between the instability regions where He and H bursts are triggered. This “stability strip” corresponds roughly to luminosities $1.3 \times 10^{36} < L < 6.5 \times 10^{36} \text{ erg s}^{-1}$. This implies that if the source is located between the Galactic Center and $\sim 16 \text{ kpc}$, the lack of bursts is to be expected, as is the lack of the detection of an IR companion.

Church et al. (2014) recently proposed a unified model for the LMXB sources. The model assumes the presence of an extended accretion disk corona (ADC). For XMM-Newton column densities are consistent with the overall Galactic Center value, it is not likely that the companion star is hidden by a cloud either. However, we cannot exclude a high-mass companion solely on the basis of the absence of a bright IR source. In fact, a B-type star placed at the distance of 8.5 kpc could have IR magnitudes fainter than the limiting H-band magnitude of the vista surveys of the Galactic Center (McMahon et al. 2013), assuming an IR absorption corresponding to the hydrogen column density we measured for 1E1743.1-2843 (see Table 2).

Since the inferred position of 1E1743.1-2843 is $\sim 16 \text{kpc}$, it is tempting to propose that the source might be a NS binary system. Consequently, this source might be the counterpart to the rare burster 1E1743.1-2843.
$L > 1 - 2 \times 10^{37}$ erg s$^{-1}$, the ADC is in thermal equilibrium with the NS surface, giving rise to a Comptonized spectrum with $E_{\text{cut}} \sim 6$ keV, which corresponds roughly to three times the actual temperature of the electrons (in the assumption of high optical depth). For lower luminosities, Comptonization becomes inefficient in cooling the corona, the thermal equilibrium assumption breaks down, and, as a consequence, the extended ADC heats up to several tens of keV. In this scenario, the ~1 km blackbody radii in LMXBs are explained by an emitting region in the shape of an equatorial strip in the orbital plane with a half height of $h \sim 100$ m. This scenario is described by a blackbody plus a cutoff power law (our model 1). The authors also note that compTT is not consistent with the evidence for an extended corona. The outer regions of the accretion disk ($kT \sim 0.1$ keV), which provide the seed photons for Compton scattering, are not expected to be detected. Our results for the blackbody temperature ($kT \sim 1.8$ keV) and radius ($r \sim 1$ km), for the $E_{\text{cut}} \sim 3kT_c$, $\sim 6.6$ keV of the Comptonized component, and for the ratio of total to blackbody luminosity for 1E1743.1-2843 fit within the expected ranges for the scenario they depict as the banana state of the Atoll sources. This interpretation, however, implies $L > 2 \times 10^{37}$ erg s$^{-1}$ to explain the spectral cutoff value and the lack of bursting activity. This would require a distance of a few tens of kiloparsecs to account for the observed flux, which could be explained if 1E1743.1-2843 were in the Sagittarius dwarf elliptical galaxy (SAGDEG), one of the small dwarf spheroidal galaxies that orbit the Milky Way. SAGDEG is currently behind the GC at a distance of $d \sim 26$ kpc (Cole et al. 2009), and is in an advanced state of destruction due to tidal interactions with the Galaxy. Therefore, a fraction of the stars that composed this dwarf galaxy have likely scattered to even greater distances.

In the HMXB scenario, the presence of a strong magnetic field ($B \sim 10^{12-13}$ G) suppresses the propagation of the bursts across the NS surface (Gilfanov & Sunyaev 2014). However, the lack of eclipses, an Fe K$\alpha$ line, and pulsations in the light curve, as well as the missing detection of a companion star, make the HMXB hypothesis less favored, even though the value of the spectral cutoff ($9 \pm 1.8$ keV) is compatible with that expected from HMXBs ($10-20$ keV).

We conclude that while an HMXB framework leaves several unexplained features, interpreting 1E1743.1-2843 as an NS-LMXB scenario is more consistent, implying a peculiar but not unique object. In this case, the source could be located at a distance of $9 < d < 16$ kpc, between the two instability luminosity intervals where He and H bursts are triggered, or, if we rely on the unified model proposed by Church et al. (2014), at a distance of $d > 36$ kpc, corresponding to the stable burning of accreted material.

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