A PILOT FOR A VLA HI DEEP FIELD

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ABSTRACT

High-resolution 21-cm HI deep fields provide spatially and kinematically resolved images of neutral hydrogen at different redshifts, which are key to understanding galaxy evolution across cosmic time and testing predictions of cosmological simulations. Here we present results from a pilot for an HI deep field done with the Karl G. Jansky Very Large Array (VLA). We take advantage of the newly expanded capabilities of the telescope to probe the redshift interval 0 < z < 0.193 in one observation. We observe the COSMOS field for 50 hours, which contains 413 galaxies with optical spectroscopic redshifts in the imaged field of 34′ x 34′ and the observed redshift interval. We have detected neutral hydrogen gas in 33 galaxies in different environments spanning the probed redshift range, including three without a previously known spectroscopic redshift. The detections have a range of HI and stellar masses, indicating the diversity of galaxies we are probing. We discuss the observations, data reduction, results and highlight interesting detections. We find that the VLA’s B-array is the ideal configuration for HI deep fields since its long spacings mitigate RFI. This pilot shows that the VLA is ready to carry out such a survey, and serves as a test for future HI deep fields planned with other SKA pathfinders.

1. INTRODUCTION

Neutral hydrogen (HI) in emission has been extensively used to probe the internal properties of the galaxies, their halos and environment in the local Universe. Surveys such as ALFALFA and HIPASS done with single-dish telescopes have detected large samples of galaxies and measured the HI content as a function of environment, size and morphology out to z ~ 0.08 (Giovanelli et al. 2005; Zwaan et al. 2003). Both interferometers and single-dish telescopes have detected HI halo gas around the Milky Way and nearby galaxies (Sancisi et al. 2005; Zwaan et al. 2003). Both interferometers and single-dish telescopes have detected HI halo gas around the Milky Way and nearby galaxies (Sancisi et al. 2005; Zwaan et al. 2003). In addition, interferometers have mapped the HI morphology and kinematics of nearby galaxies in different environments. For example, work by Chung et al. (2009) has shown that spiral galaxies get stripped of their HI gas due to the hot gas present in the Virgo Cluster. On the other hand, galaxies residing in voids tend to have small optical disks with large HI disks, with some showing signs of on-going accretion (Kreckel et al. 2012).

Until recently, the HI Universe in emission at z > 0.1 was virtually unknown due to limited instantaneous frequency coverage and sensitivity, except for a few studies. These include the first HI detection at z > 0.1 reported by Zwaan et al. (2001) using the Westerbork Synthesis Radio Telescope (WSRT), and Arecibo observations at z = 0.1 – 0.2 (Catinella et al. 2008; Freundling et al. 2011). Indirect detections have also been reported, such as the stacked HI signal at z = 0.24 and z = 0.37 by Lah et al. (2007, 2009), and intensity mapping at z = 0.8 (Chang et al. 2010; Masui et al. 2013).

Current technology now allows for much wider instantaneous velocity coverage. The WSRT observations of two clusters in the interval 0.164 < z < 0.224 are the first ones to probe large volumes at higher redshift, detecting over 160 galaxies in HI in different environments (Verheijen et al. 2007; Jaffé et al. 2012). SKA pathfinders will be able to cover instantaneously HI from z = 0 to z = 0.23 (Apertif), z = 0.4 (ASKAP; Johnston et al. 2008), and z ~ 1 (MeerKat, Holwerda et al. 2012). The recently upgraded Karl G. Jansky Very Large Array (VLA) can now observe HI from z = 0 to z = 0.45 in one setting. Here we present results from a pilot for such a survey done during commissioning, when it was only possible to observe the redshift interval 0 < z < 0.193. The goal of this pilot is to test the fea-
we separated the 32 spectral windows into individual uv datasets. During this process, we also applied Hanning smoothing to remove the Gibbs ringing phenomenon due to strong RFI. After this, we did a preliminary calibration following standard procedures in order to be able to run RFLAG, an automatic RFI flagging task of AIPS. Following this, a final calibration was performed and applied to the target source. The calibrated target source data for each session were further flagged with RFLAG, and then the 10 sessions of every spectral window were combined using STUFFR, an AIPS procedure that allows for baseline-length-dependent time averaging. We ran this procedure by setting parameters that would not distort the sources near the edges of the field of view due to time averaging. The final uv datasets were a factor of three smaller than the combined 10 sessions without the above mentioned averaging, speeding up the imaging.

We made total intensity image cubes of $34' \times 34'$ covering the full width half power of the primary beam, using a weighting scheme intermediate between natural and uniform. Continuum subtraction was performed in the image plane by fitting Chebychev polynomials whereby data points deviating more than $2\sigma$ from the fit were iteratively removed.

## 3. Results

### 3.1. Noise and RFI

We first analyze the RFI environment and how it affected our observations since this is one of the main challenges when carrying out an HI deep field. There were three kinds of RFI present in the data: broad and strong RFI due to GLONASS and GPS, narrow and strong spikes mostly from the Albuquerque airport radar, and broad weaker signals of unknown origin; none of which varied over timescales shorter than minutes.

The top panel of Figure 1 shows that we reach the theoretical noise of 0.2 mJy beam$^{-1}$ in most of the observed frequencies. We find that the noise is higher in the frequency interval $1225 - 193$ MHz where strong RFI was affected our observations since this is one of the main challenges when carrying out an HI deep field. There were three kinds of RFI present in the data: broad and strong RFI due to GLONASS and GPS, narrow and strong spikes mostly from the Albuquerque airport radar, and broad weaker signals of unknown origin; none of which varied over timescales shorter than minutes.

The top panel of Figure 1 shows that we reach the theoretical noise of 0.2 mJy beam$^{-1}$ in most of the observed frequencies. We find that the noise is higher in the frequency interval $1225 - 193$ MHz where strong RFI was present due to GLONASS and GPS, and a high percentage of the data was flagged. The effect of flagged data can also be seen in the middle panel, where the beam size is smoothly decreasing with increasing frequency (as expected), with some sharp drops at the frequencies affected by RFI. This is due to most of the short spacings getting flagged, thus making a smaller synthesized beam.
is a major issue, compared to ASKAP, MeerKAT, and Apertif, for which most of the baselines are considerably shorter than 5 km.

3.2. Source Finding

We apply additional frequency averaging on the continuum subtracted image cubes to increase the signal to noise for source finding, resulting in an effective frequency resolution of 125 kHz (26.8 km s\(^{-1}\) at z = 0) and a noise per channel of 0.085 mJy beam\(^{-1}\). We have done an extensive search by eye to have a reliable sample of detections to compare with the results from automated source finding algorithms. A blind search only yields the brightest detections, prompting us to do a more targeted search for HI associated with each of the 413 galaxies with previously known spectroscopic redshift. All of the detections have spatially coherent emission seen in at least three consecutive channels with at least 3\(\sigma\) in two of them in the immediate vicinity of the galaxy, within 1 MHz (~200 km s\(^{-1}\)) of the frequency corresponding to the optical redshift.

The search by eye yields 33 detections, with three lacking a measured optical spectroscopic redshift. This number of detections is roughly consistent with predictions based on photometric HI masses (Kannappan 2004, Catinella et al. 2010). The optical spectroscopic sample is not complete, especially for small systems and low redshifts such as the three mentioned above. Some of the total HI images at z > 0.1 appear to be incomplete based on the asymmetric distribution and kinematics, suggesting we are only seeing the highest column density gas.

We are in the process of testing several automated source finding algorithms as compared in Popping et al. (2012), with an emphasis on DUCHAMP and S+C FINDER (Whiting 2012, Serra et al. 2012). The searches are being done both blindly and using the spatial and frequency information from the optical catalogues. We will present the full data catalogue plus a detailed analysis of the detections and detection methods in a follow-up paper.

3.3. Measuring the HI Properties

There are conflicting equations in the literature on how to calculate HI mass and column density at cosmologically significant distances due to different velocity definitions. Here we show how to calculate these quantities in terms of frequency. We start with the fundamental relation between the observed flux S and luminosity L:

\[
L_{HI} = 4\pi D_L^2 S_{HI}
\]

with \(D_L\) being the luminosity distance, and \(S_{HI} = \int s(\nu) \, d\nu\) where \(s(\nu)\) is the observed flux density and \(d\nu\) is the frequency width in the observer’s frame. The luminosity can be converted to HI mass using atomic physics and assuming the line is optically thin, resulting in the following equation:

\[
M_{HI} = 49.8 \, D_L^2 \int s \, d\nu \, [M_\odot]
\]

where \(D_L\) is in Mpc and the integral is in units of Jy Hz. We can calculate the column density by dividing the
Fig. 2.— Optical spectroscopic sample and our detections in HI. **Top:** Three-dimensional distribution of galaxies in the data cube with the coordinates translated to a linear distance from the pointing center. The gray circles represent the galaxies with known optical spectroscopic redshifts, and the red symbols show the HI detections. The circles are detections with known spectroscopic redshifts, while the squares show our three detections without one, where we calculate the redshift from the HI spectrum. **Bottom:** Projection of 3D plot in RA-z with HI distribution maps highlighting the detections. HI emission contours are overlaid on COSMOS HST images (Koekemoer et al. 2007; Massey et al. 2010), with the synthesized beam plotted in the bottom-right corner of each panel. Starting from the bottom-left and moving counterclockwise: an isolated galaxy, a galaxy pair, an on-going merger in a wall, and the highest redshift detection. Following the same order, the contours correspond to the following HI column density values: (1.8, 3.6, 7.2, 14.4)×10^{20}, (2.1, 4.2, 6.3, 8.4)×10^{20}, (2.7, 5.4, 8.1, 10.8)×10^{20}, and (3.5, 7.0, 10.5, 14)×10^{20} atoms cm^{-2}. The opening angle of the pie-diagram is greatly exaggerated for clarity.
HI mass by the physical size covered by the beam. The derived equation is the following:

\[ N_{HI} = \frac{2.34 \times 10^{20}}{\theta_1 \theta_2} (1 + z)^4 \int s \, dv \, [\text{cm}^{-2}] \]  

(3)

where \( \theta_1 \) and \( \theta_2 \) are the FWHM sizes of the major and minor axes of the synthesized beam in arcseconds, and \( z \) is the redshift.

3.4. Detections in Different Environments and Redshifts

The top panel of Figure 2 shows that the distribution of galaxies in the survey volume is not uniform; instead we are probing different environments from nearby galaxies in voids to galaxies that reside in walls. There are two over-dense regions: a nicely defined wall at \( z = 0.12 \) and a larger concentration of galaxies in the interval \( 0.17 < z < 0.19 \). There are empty regions, such as at \( z < 0.04 \), and a very noticeable gap from \( z = 0.13 \) to \( z = 0.16 \), where only four galaxies are present. The detections, as shown by the red symbols, span the probed redshift range, with most of them residing in the above mentioned wall that has over 100 galaxies with known spectroscopic redshift.

The bottom panel of Figure 2 highlights some of the detections and their location on a RA-\( z \) projection of the surveyed volume. The first galaxy shown (bottom left) resides in an under-dense region, and is one of the detections that did not have a measured optical redshift. The HI is distributed in a disk that is almost three times bigger than the optical one, and is asymmetric, with the HI emission being more extended toward the north. We also detect galaxies with companions, an example of this is the pair of galaxies detected at \( z = 0.076 \) (bottom right). The larger galaxy on the left has an HI distribution that consists of two components, with the southern component showing an extension to the right in the direction of a companion that is 83 kpc away, and is also detected in HI. We also highlight one of the detections in the wall at \( z = 0.12 \) (top right), with the optical morphology showing several stellar components suggesting it is an on-going merger. The highest redshift detection is shown in the top-left corner, where the HI is offset from the optical image. The two smaller galaxies shown in the panel do not have optical spectroscopic redshifts.

In a follow-up paper, we will present our detections and a comprehensive study of the HI properties with respect to the large scale structure, here we present some preliminary analysis on the properties of our detections in Figure 3 in terms of the HI mass, stellar mass, and \( NUV - r \) color. The colors were attained from GALEX and SDSS DR8, and have been corrected for extinction and band-shifting (k-correction) effects. We were able to get reliable photometry for 28 of the detections, the remaining five were either not detected by GALEX or were in close proximity to a star. The stellar masses are calculated using SDSS colors and the K band luminosity following Bell et al. (2003).

As seen in the top left panel, we have detected a range of HI masses, from \( 1.6 \times 10^7 M_\odot \) to \( 1.4 \times 10^{10} M_\odot \), consistent with the predicted sensitivity at different redshifts assuming a 150 km s\(^{-1}\) profile width. The masses below the sensitivity limit are either associated with galaxies that have a smaller profile width, or they were found in the search by eye at the position and redshift of a spectroscopically known galaxy, where we used a different detection criterion (see section 3.2). The color-magnitude diagram (top right) shows that the majority of the galaxies within our observed volume are blue. The HI detections mostly correspond to galaxies that are star-forming, as indicated by the fact that we only detect two galaxies with \( NUV - r > 4 \). The third panel shows that the galaxies occupy a relatively narrow range in HI mass compared to the stellar mass, with the smallest galaxies being gas dominated (e.g., Kreckel et al. 2012). We find consistency when comparing to ALFALFA (Huang et al. 2012), especially for high stellar masses. We also compare the measured gas fractions and color with work by Catinella et al. (2012), and find agreement between the two surveys.

3.5. Stacking

We performed stacking of the HI spectra of the 80 galaxies with known optical zCOSMOS redshifts between 0.12 and 0.13, corresponding to the large-scale structure visible in Figure 2, to illustrate the potential of the data for statistical studies. We find an average HI mass of \( 1.8 \pm 0.3 \times 10^9 M_\odot \). This is somewhat smaller than other stacking studies find at this redshift (e.g. Delhaize et al. 2013 submitted, Gereb et al. in prep). This is most likely due to the fact that the galaxies at this redshift in the zCOSMOS survey tend to be somewhat smaller than used in these other studies.

4. SUMMARY AND CONCLUSIONS

We have, for the first time, carried out an HI deep field probing the entire redshift range \( 0 < z < 0.193 \). The pilot data give a glimpse of what will be possible with an HI deep field. With only 50 hours, we identify 33 detections, 3 of them without an optical spectroscopic redshift. These data will allow us to study the evolution of HI properties for a number of galaxies as a function of environment using high-resolution HI maps and stacking techniques.

Our pilot has shown that the VLA is ready to carry out an HI deep field. We have successfully used automatic flagging algorithms to remove most of the RFI and were able to get useful data across the 1190-1426 MHz band. The impact of RFI was mostly seen at spacings \( < 2 \) km, where as much as 30-40% of the data were flagged, and baselines longer than 4 km are largely unaffected by RFI. This argues in favor of using configurations with most of the collecting area at spacings longer than 4 km, even though this may result in slightly reduced surface brightness sensitivity. The VLA’s B-array, where 70% of the baselines are longer than 2 km, is the ideal configuration for an HI deep field.

We are entering a new era in radio astronomy and the results of the HI deep fields planned with the VLA and other SKA pathfinders will increase our understanding of galaxy evolution. Assuming the same parameters as in the pilot, a 1000 hour observation done with the VLA of the same pointing over the redshift range \( 0 < z < 0.45 \), will translate to approximately 300 5σ detections. These observations will place better constraints on simulations, providing them with HI properties in a wide range of stellar masses and environments across cosmic time.
Fig. 3.— Properties of the detections with the colors corresponding to redshift and the transparent symbols showing those that lie below the 5σ solid curve. The symbols correspond to the following: galaxies with known optical redshift (circles), galaxies with unknown optical redshift (squares), and galaxies with $NUV - r > 4$ (triangles). Top left: HI mass as a function of redshift for the 33 detections, with the 5σ sensitivity curve assuming 150 km s$^{-1}$ (solid black line) and 50 km s$^{-1}$ (dashed line). Top right: Color magnitude diagram for the galaxies within the search volume shown in black, with the color symbols showing the detections. Bottom left: HI mass as a function of stellar mass with the dashed lines showing the ALFALFA values [Huang et al. 2012]. Bottom right: Ratio of these two quantities as a function of $NUV - r$ as compared to [Catinella et al. 2012].

Fig. 4.— Stacked signal of 80 galaxies with optical redshifts from ZCOSMOS in the interval $0.12 < z < 0.13$. The average HI mass is $1.8 \pm 0.3 \times 10^9 M_\odot$

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REFERENCES

Bell, E. F., McIntosh, D. H., Katz, N., & Weinberg, M. D. 2003, ApJS, 149, 289

Catinella, B., Haynes, M. P., Giovanelli, R., Gardner, J. P., & Connolly, A. J. 2008, ApJ, 685, L13
